Borexino results after 200 days of data

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(on behalf of the Borexino collaboration)

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Borexino Collaboration

Milano

Genova

Perugia

Princeton University

Kurchatov Institute (Russia)

Heidelberg (Germany)

Virginia Tech. University

Jagiellonian U. Cracow (Poland)

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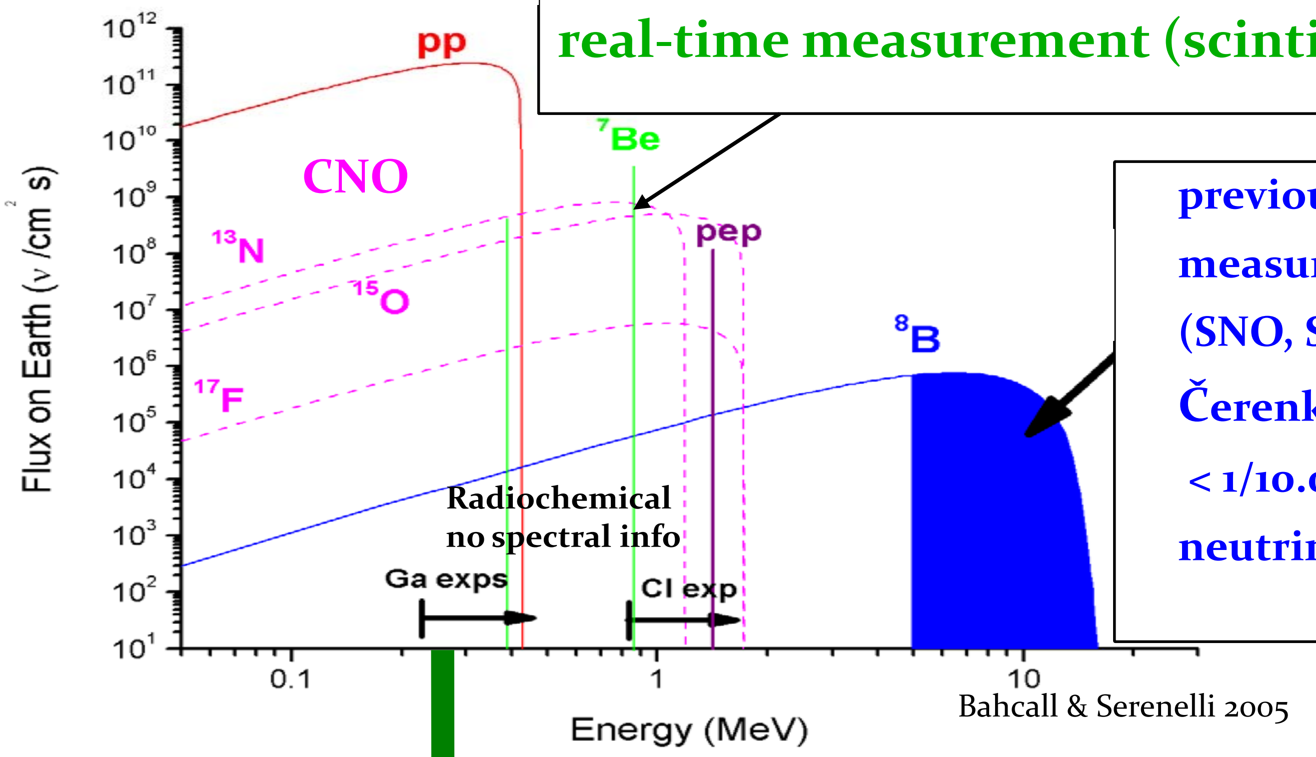
Munich (Germany)

APC Paris

Solar neutrino energy spectrum

mono-energetic 0.862 MeV ⁷Be-ν

real-time measurement (scintillation light)



previous real-time measurements (SNO, SuperKamiokande) Čerenkov radiation < 1/10.000 of the total solar neutrino flux

Borexino energy threshold

Standard Solar Model: Neutrino fluxes vs solar metallicity

(metallicity – abundance of the elements above Helium)

Φ ($cm^{-2}s^{-1}$)	pp ($x\ 10^{10}$)	pep ($x\ 10^{8}$)	^{7}Be ($x\ 10^{9}$)	^{8}B ($x\ 10^{6}$)	^{13}N:CNO ($x\ 10^{8}$)	^{15}O:CNO ($x\ 10^{8}$)	^{17}F:CNO ($x\ 10^{6}$)
BS05 [1] GS 98 [2]	5.99	1.42	4.84	5.69	3.07	2.33	5.84
BS05 [1] AGS05[3]	6.06	1.45	4.34	4.51	2.01	1.45	3.25
Δ	+1%	+2%	-10%	-21%	-35%	-38%	-44%

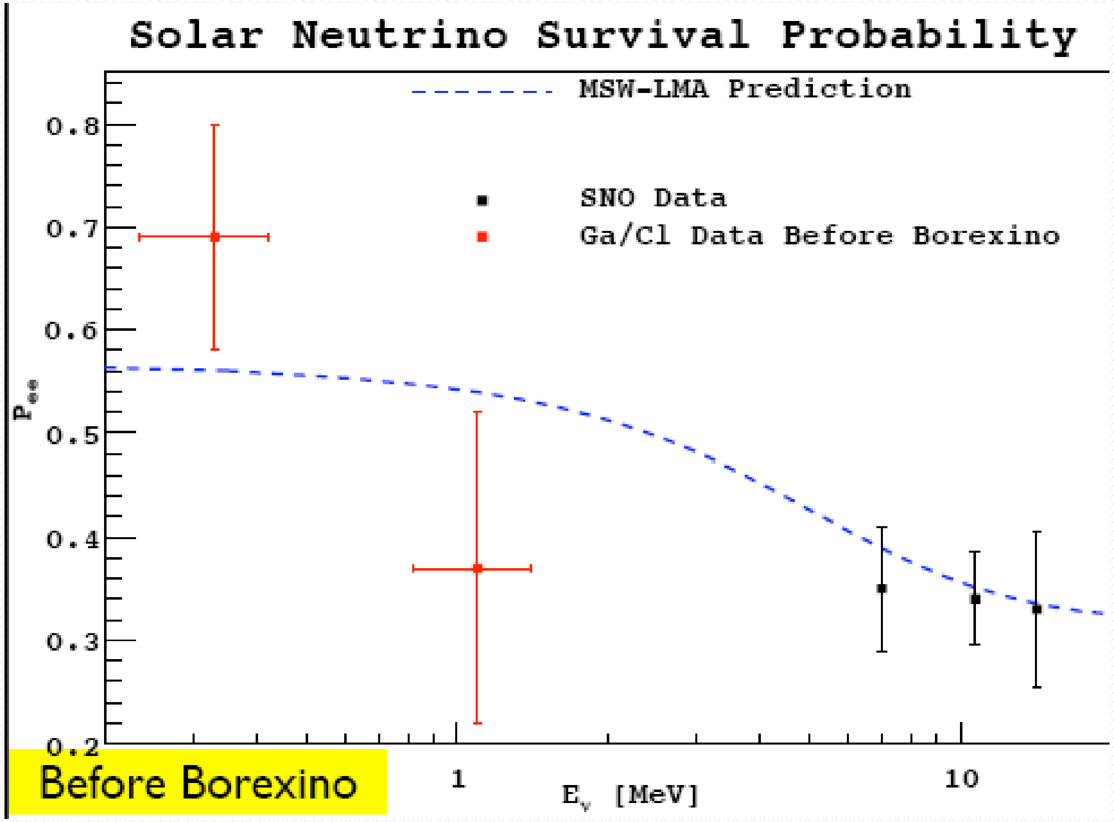
[1]BS05: Bahcall, Serenelli & Basu, AstropJ 621 (2005) L85

[2]Based on **high metalicity model** GS98: Grevesse & Sauval, Space Sci. Rev. **85, 161 (1998)**

[3]Based on new **low metalicity model** AGS05 (factor ~2):
Asplund ,Grevesse & Sauval 2005, Nucl. Phys. A **777, 1** (2006).
BUT: incompatible with helioseismological measurements

MEASURING for the first time the CNO-neutrino fluxes would help to resolve the controversy!

Solar –ν survival probability BEFORE BOREXINO



Vacuum regime

Matter regime

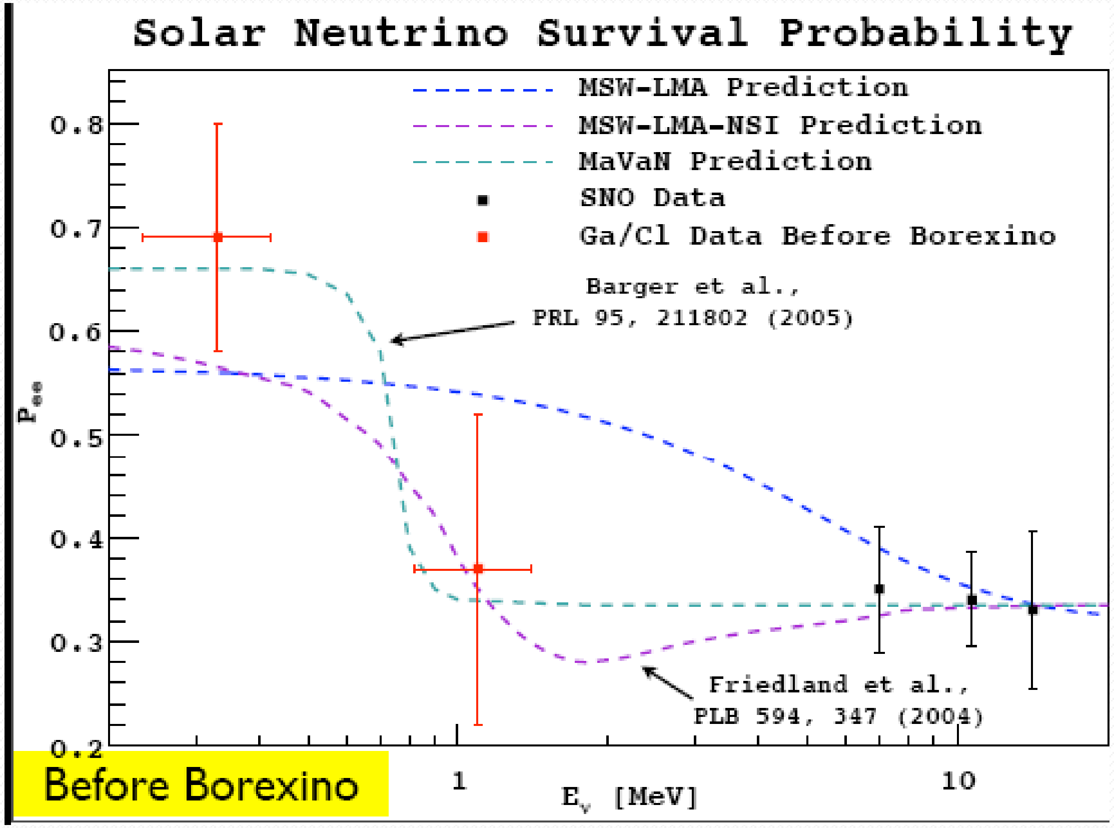
Transition region of 1-2 MeV:
First time observed by Borexino

Low energy neutrinos (pp):
flavor change dominated
by vacuum oscillations;

High energy neutrinos (^8B):
Resonant oscillations in matter
(MSW effect):
Effective electron neutrino mass
is increased due to the charge
current interactions
with electrons of the Sun

Transition region (^7Be):
Decrease of the ν_e survival
probability (P_{ee})

Solar –ν survival probability BEFORE BOREXINO



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Transition region of 1-2 MeV:
First time observed by Borexino

Scientific goals of Borexino

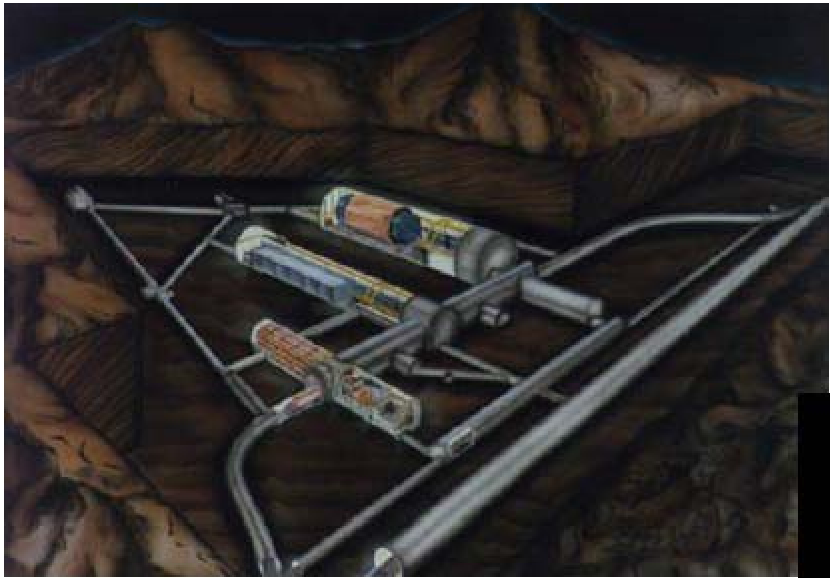
- **The first real-time measurement of sub-MeV solar neutrinos;**
- The first simultaneous measurement of solar neutrinos from the transition region (**^{7}Be-ν – two measurements published**) and from the matter-enhanced oscillation region (**^{8}B-ν – the first measurement below 5 MeV submitted**)
- Precision measurement (at or below the level of 5%) of the ^{7}Be-ν rate: to test the SSM and MSW-LMA solution of the Standard Solar Problem as well as the balance between the neutrino and photon luminosity of the Sun;
- Check the 7% **seasonal variation** of the neutrino flux (confirm solar origin);
- Under study: first measurement of the **CNO** neutrinos (sun metallicity controversy);
- Under study: ***pep*** neutrinos – indirect constrain on the pp-flux;
- High energy tail of ***pp*** neutrinos;
- Antineutrinos and **geoneutrinos**;
- **Supernovae** neutrinos and antineutrinos;

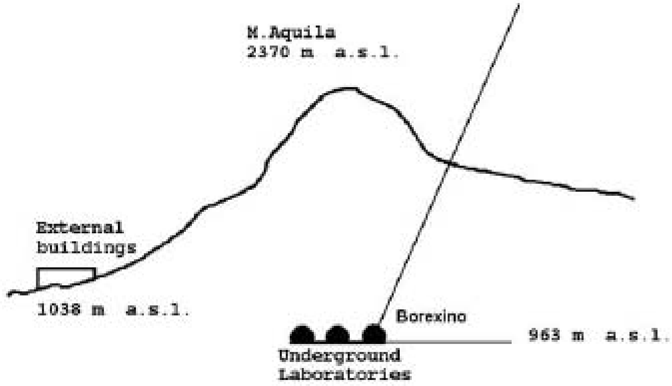
Detection principles in Borexino

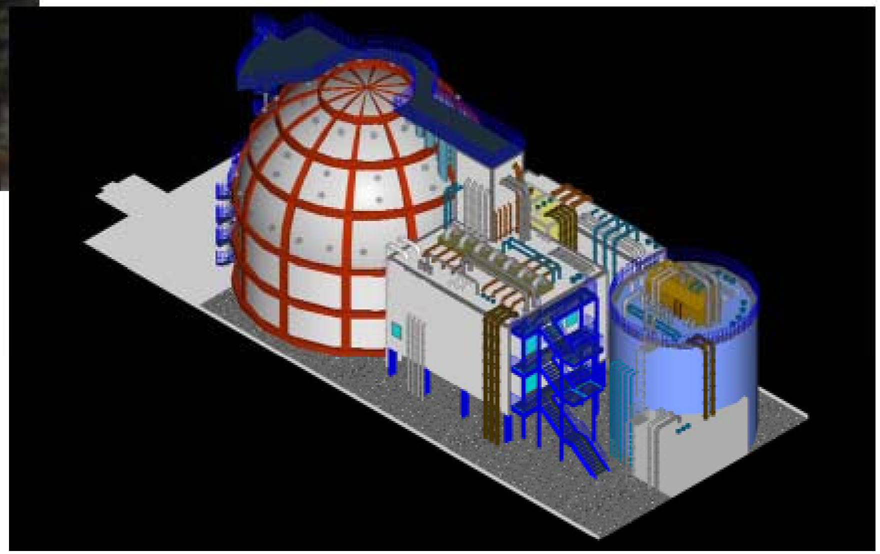
- Neutrino elastic scattering on electrons of liquid scintillator;
- Scattered electrons cause the scintillation light production;
- Advantages:
  - Low energy threshold (0.2 MeV);
  - Good energy resolution;
  - Good position reconstruction;
- Drawbacks
  - Info about the neutrino directionality is lost ;
  - ν-induced events can't be distinguished from the events of β/γ natural radioactivity;

Extreme radiopurity is a must!

Experimental site







Borexino is located at the **Laboratori Nazionali del Gran Sasso**, near L'Aquila, shielded by 1400 m of Rocks (~3500 m water equivalent)

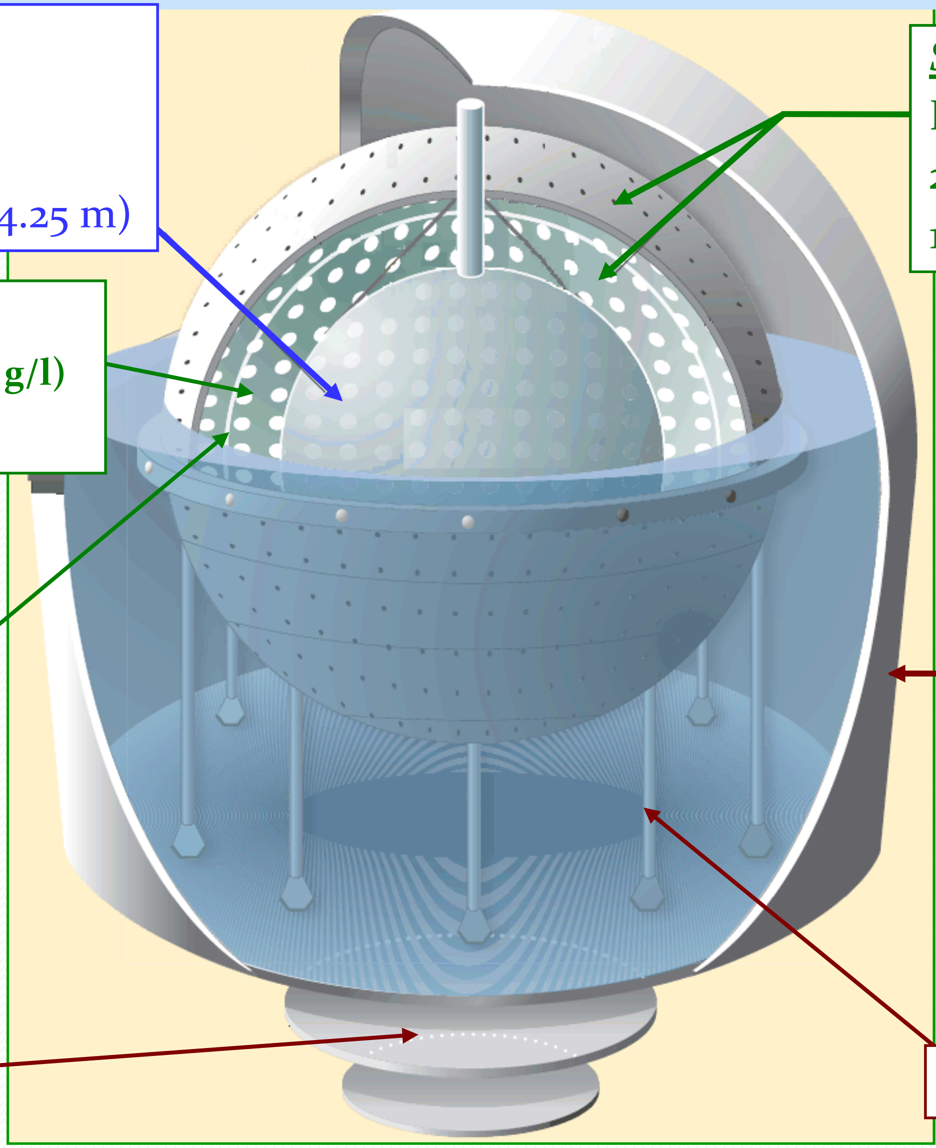
Borexino Detector

Scintillator:
270 t PC+PPO (1.5 g/l)
in a 150 μm thick
inner nylon vessel (R = 4.25 m)

Buffer region:
PC+DMP quencher (5 g/l)
4.25 m < R < 6.75 m

Outer nylon vessel:
R = 5.50 m
(^{222}Rn barrier)

Carbon steel plates



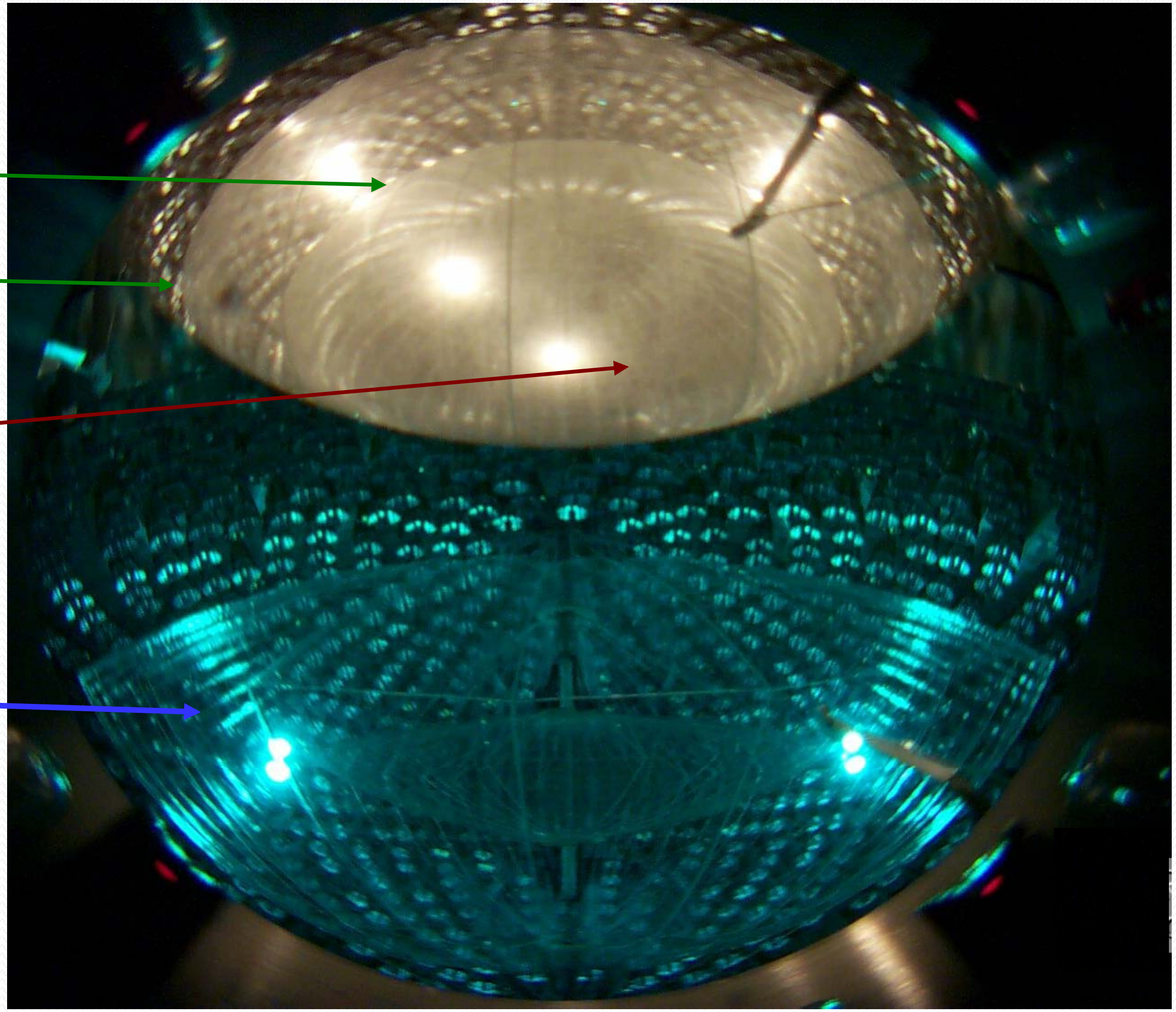
Stainless Steel Sphere:
R = 6.75 m
2212 PMTs
1350 m^3

Water Tank:
γ and n shield
μ water Č detector
208 PMTs in water
2100 m^3

20 steel legs

Filling 1/3

End October 2006



Nylon Vessels
Inner: 8.5 m
Outer: 11.0 m

LAKN – Low Argon and Krypton Nitrogen

Ultra-pure water

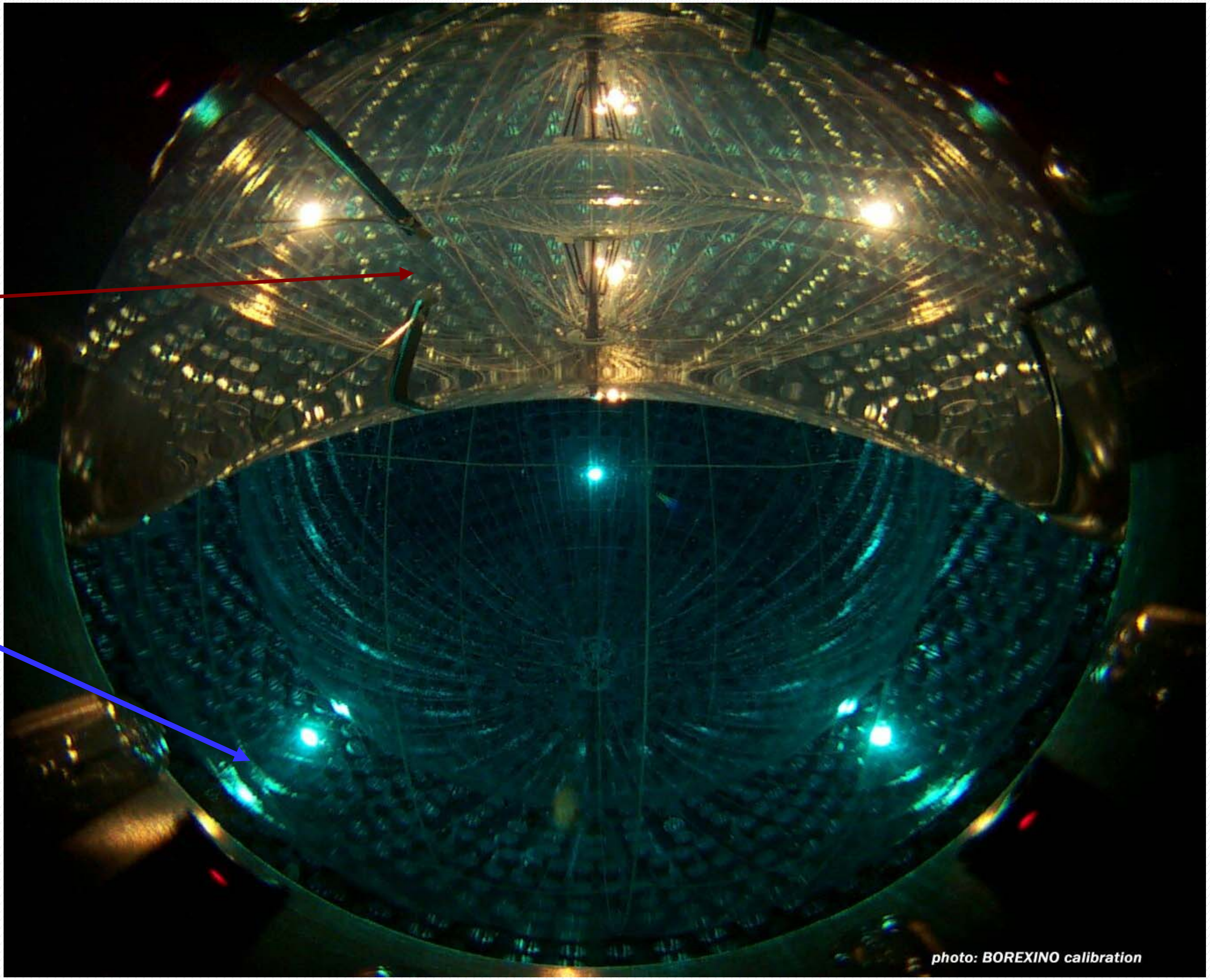
Foto taken with one of 7 CCD cameras placed inside the detector

Filling 2/3

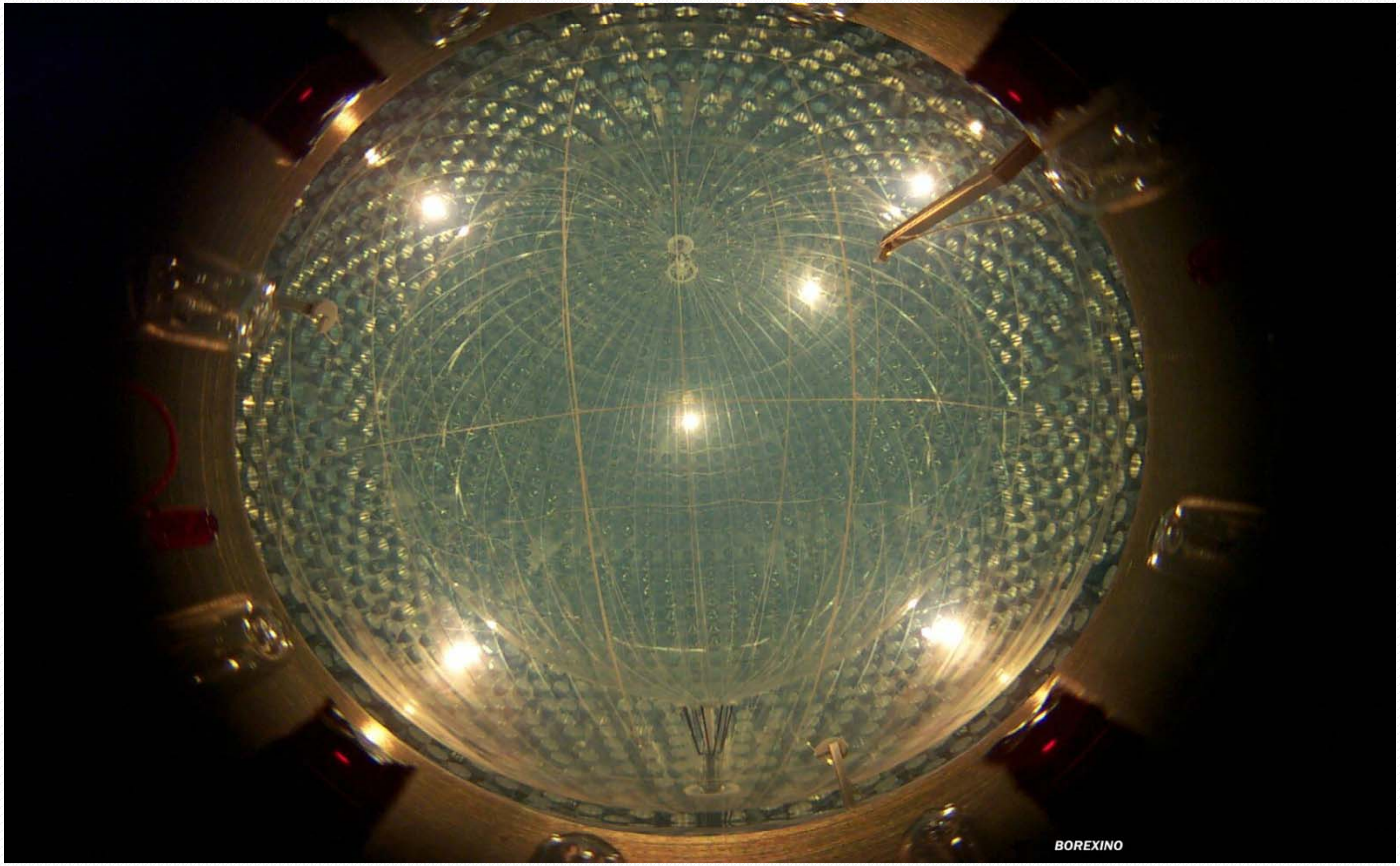
March 2007

Liquid scintillator

Ultra-pure water

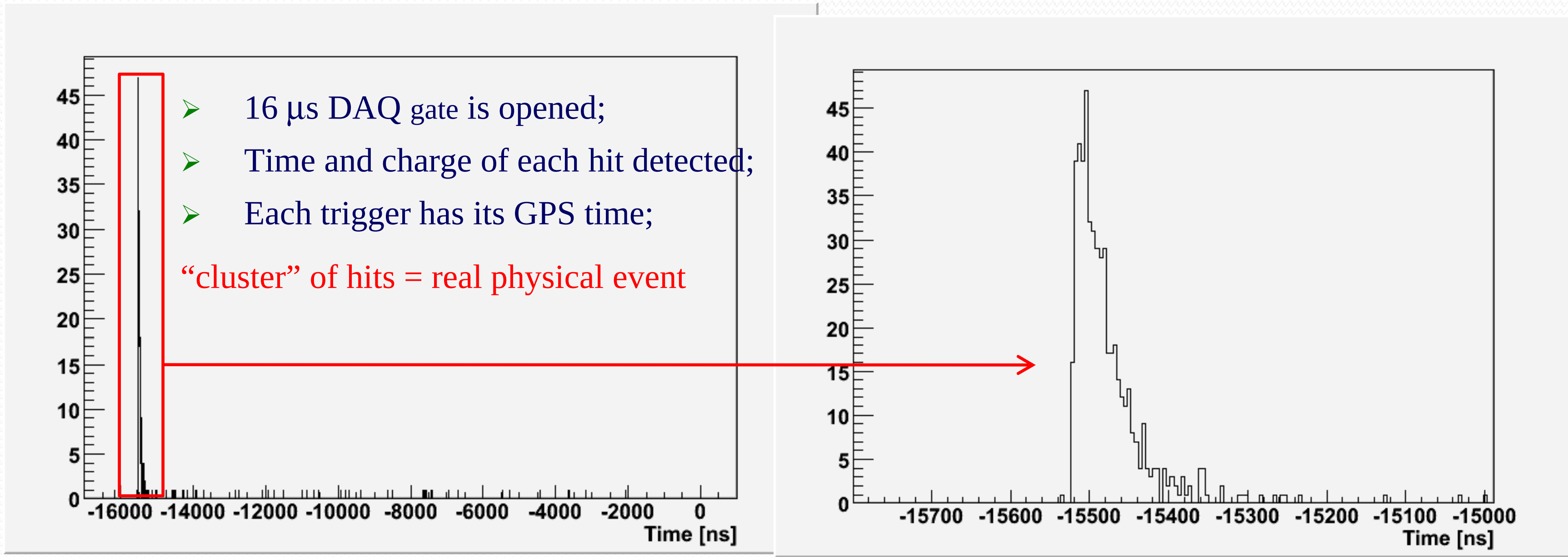


Detector filling completed on May 15th, 2007



Data acquisition and data structure

- Charged particles and γ produce scintillation light: photons hit inner PMTs;
- DAQ trigger: > 25 inner PMTs (from 2212) are hit within 60-95 ns:



- Outer detector gives a muon veto if at least 6 outer PMTs (from 208) fire;

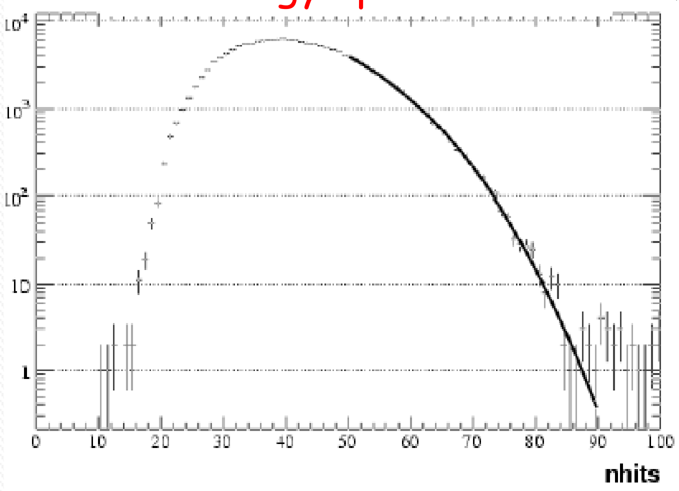
Detector performances

Lighth yield

1) ^{14}C spectrum, β^- decay 156 keV end point;
2) ^{11}C spectrum, β^+ decay 960 keV, triple coincidence with muon and neutron;
3) Global spectral fit (^{14}C, ^{210}Po, ^7Be edge);

(500 ± 12) p.e./MeV

taking into account quenching factor



Energy resolution: 10% @ 200 keV
8% @ 400 keV
6% @ 1000 keV

Spatial resolution: 35 cm @ 200 keV
(scaling as $N_{p.e.}^{-1/2}$) 16 cm @ 500 keV

Fiducial volume definition

- the nominal Inner Vessel radius: 4.25 m (278.3 tons of scintillator)
- how to define fiducial volume of 100 tons?
  1) rescaling background components known to be uniformly distributed within the scintillator (^{14}C bound in scintillator itself, capture of μ-produced neutrons on protons)
  2) using the sources with known position: (Th emitted by the IV-nylon, γ external background, teflon diffusers on the IV surface)

Background: ^{232}Th and ^{238}U content

Assuming secular equilibrium:

^{232}Th chain

τ = 432.8 ns

$^{212}Bi \xrightarrow[2.25\ MeV]{\beta} {}^{212}Po \xrightarrow[\sim 800\ keV\ eq.]{\alpha} {}^{208}Pb$

^{238}U chain

τ = 236 µs

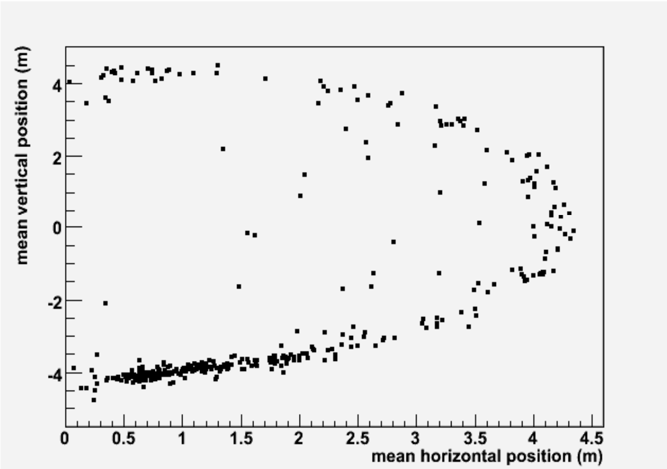
$^{214}Bi \xrightarrow[3.2\ MeV]{\beta} {}^{214}Po \xrightarrow[\sim 700\ keV\ eq.]{\alpha} {}^{210}Pb$

(6.8±1.5) × 10^{-18} g(Th)/g

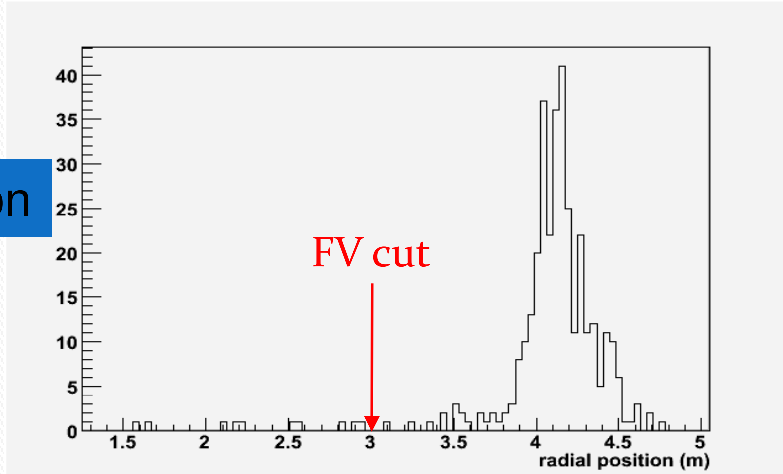
Bulk contamination

(1.6 ± 0.1) × 10^{-17} g(U)/g

Only few bulk candidates in the FV



Thoron



Background: ^{210}Po and ^{85}Kr

^{210}Po: end of ^{238}U chain :

	^{210}Pb	$\xrightarrow{\beta^-(61\ keV)}$	^{210}Bi	$\xrightarrow{\beta^-(1.2MeV)}$	^{210}Po	$\xrightarrow{\alpha}$	^{206}Pb
$t_{1/2}$	22.3 y		5.01 d		138.38 d		stable

- The bulk ^{238}U and ^{232}Th contamination is **negligible**
- The ^{210}Po background is **NOT related** neither to ^{238}U nor to ^{210}Pb contamination
- **May 2007 ~80 counts/day, τ = 204.6 days**
- **^{210}Bi no direct evidence** ---> free parameter in the total fit, cannot be disentangled, in the ^{7}Be energy range, from the CNO

^{85}Kr β-decay energy spectrum similar to the ^{7}Be recoil electron

$$^{85}Kr \xrightarrow[687\ keV]{\beta} {}^{85}Rb$$

τ = 10.76 y - BR: 99.56%

^{85}Kr is studied through :

$$^{85}Kr \xrightarrow[173\ keV]{\beta} {}^{85m}Rb \xrightarrow[514\ keV]{\gamma} {}^{85}Rb$$

τ = 1.46 ms - BR:0.43%

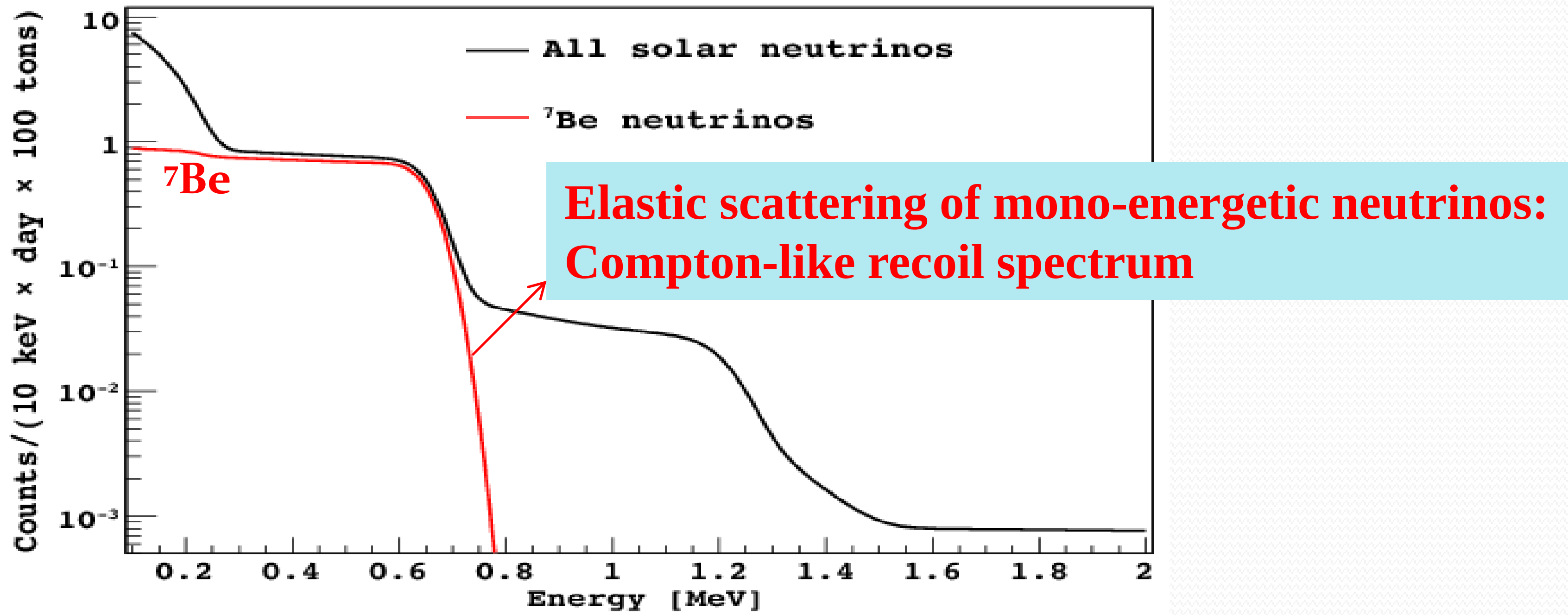
Only 8 (β–γ) coincidences selected in the inner vessel in 192 days

the ^{85}Kr contamination **(29±14) counts/day/100 ton**

More statistics is needed → taken as a free parameter in the total fit

Simulated spectrum of solar neutrinos

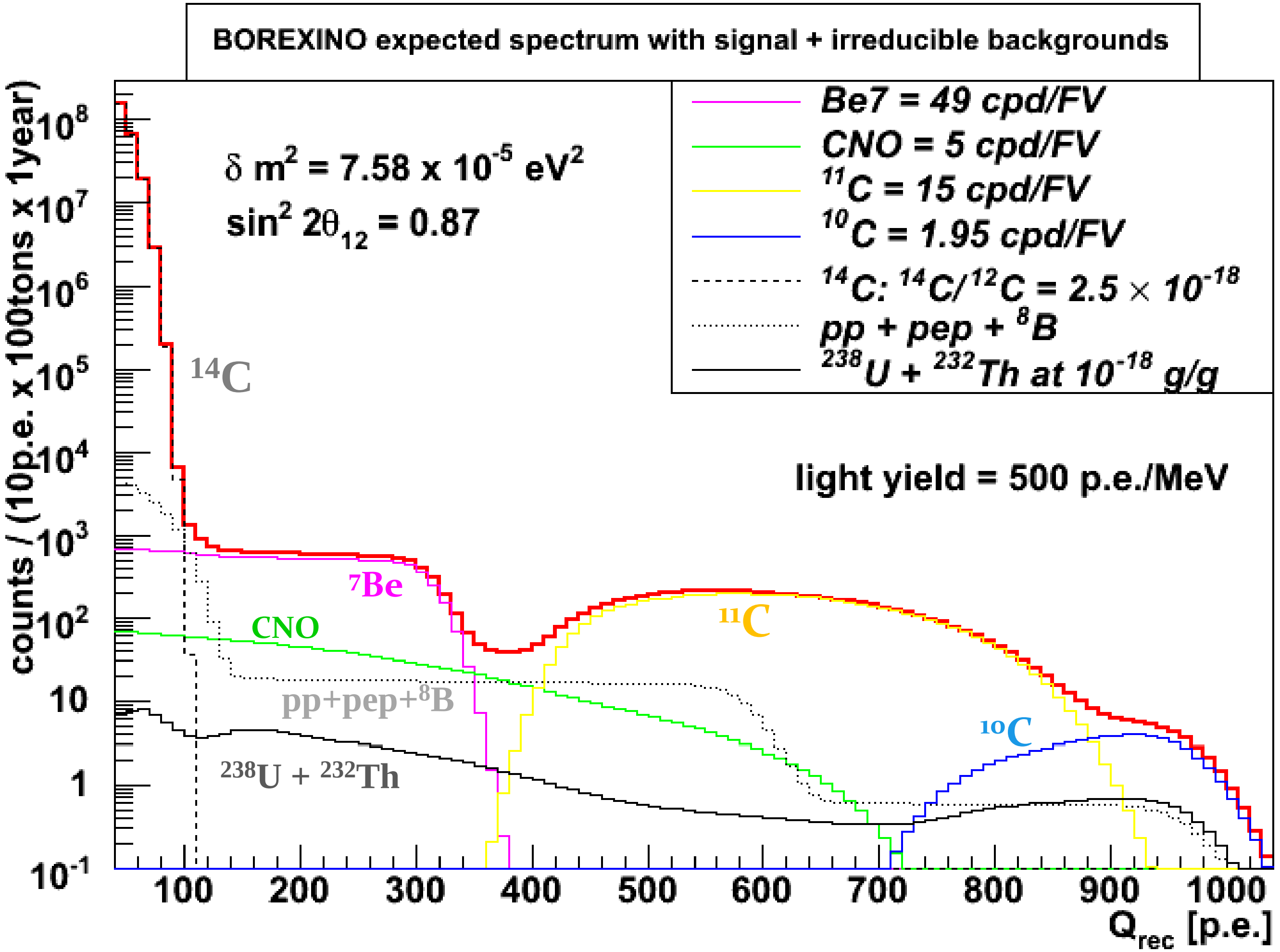
(detected via elastic scattering off electrons)



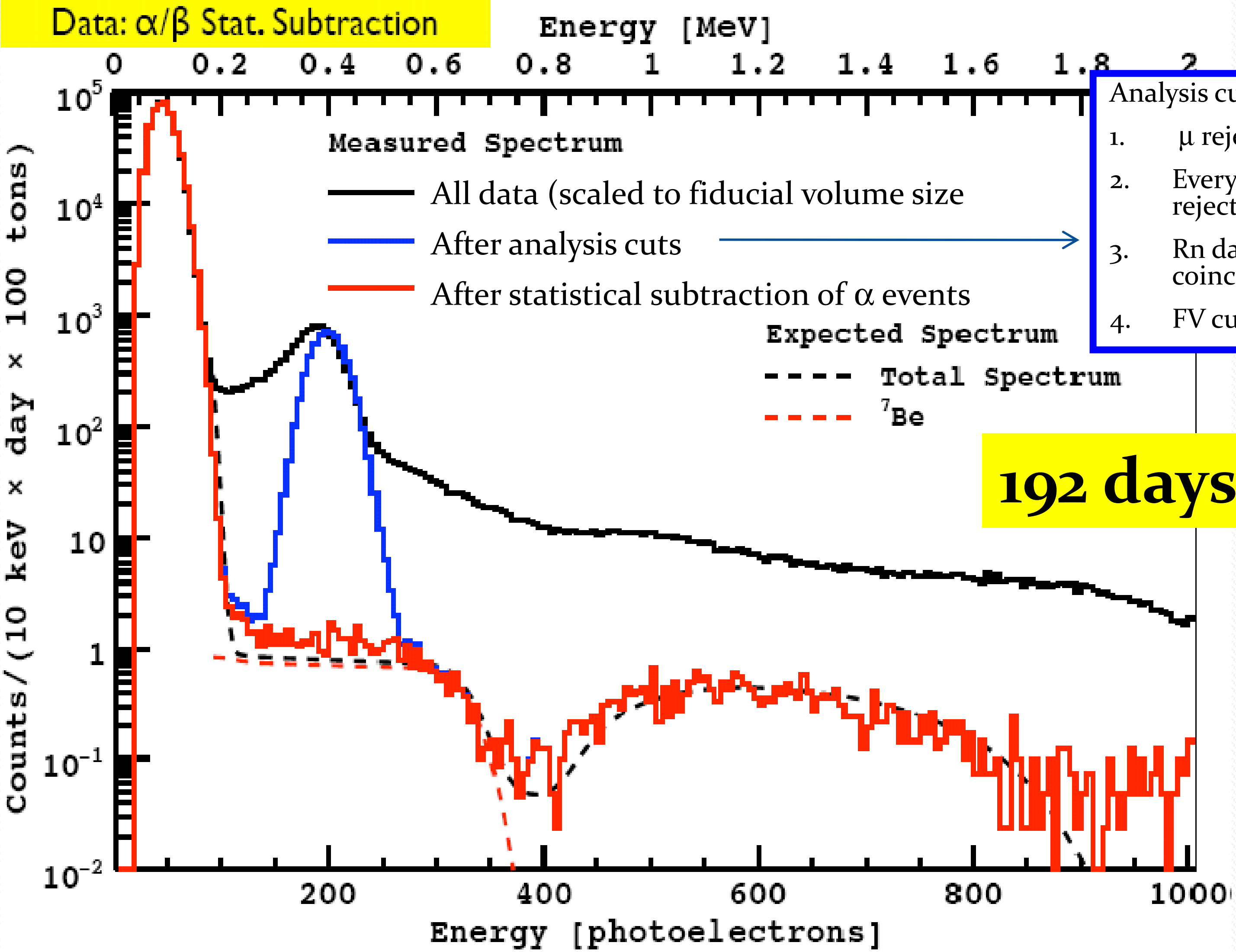
A measured neutrino flux depends on:

- neutrino flux produced in the Sun (**SSM: Standard Solar Model**)
- **neutrino survival probability**
  (LMA – Large Mixing Angle solution in the Δm^2 – $\sin^2 2\theta$ parameter space)
- **interaction cross section** (cca. 10^{-44} cm^2!)

Expected MC spectrum: signal+irreducible background



Data: α/β Stat. Subtraction

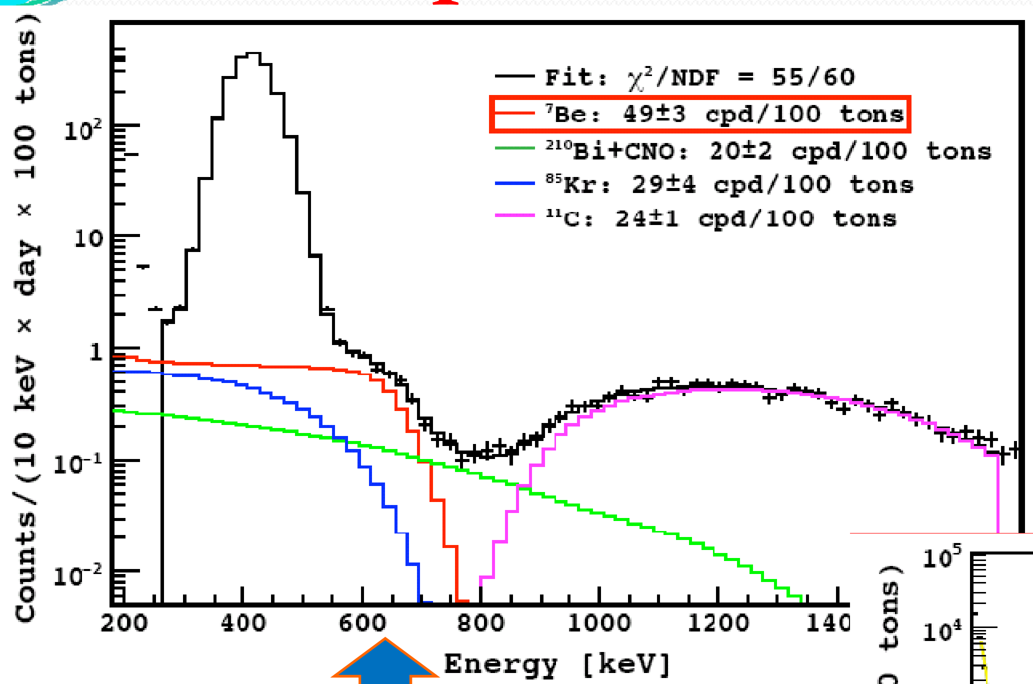


Analysis cuts

1. μ rejected;
2. Everything 2ms after μ rejected;
3. Rn daughters before Bi-Po coincidences vetoed;
4. FV cut;

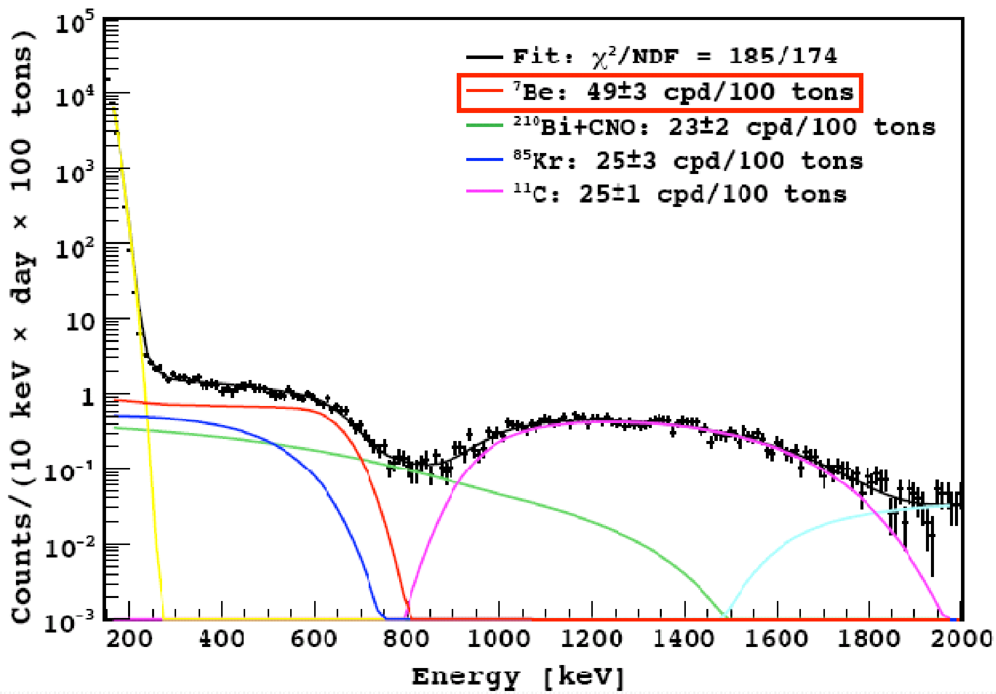
192 days of data

Spectral fit



- Fit between 100-800 p.e.;
- Light yield: a free fit parameter;
- Light quenching included (Birks' parametrization);
- ^{14}C, ^{11}C and ^{85}Kr free fit parameters;

- Fit to the spectrum **without** and **with** α–subtraction is performed giving consistent results



^{7}Be: (49 ± 3) cpd/100 tons

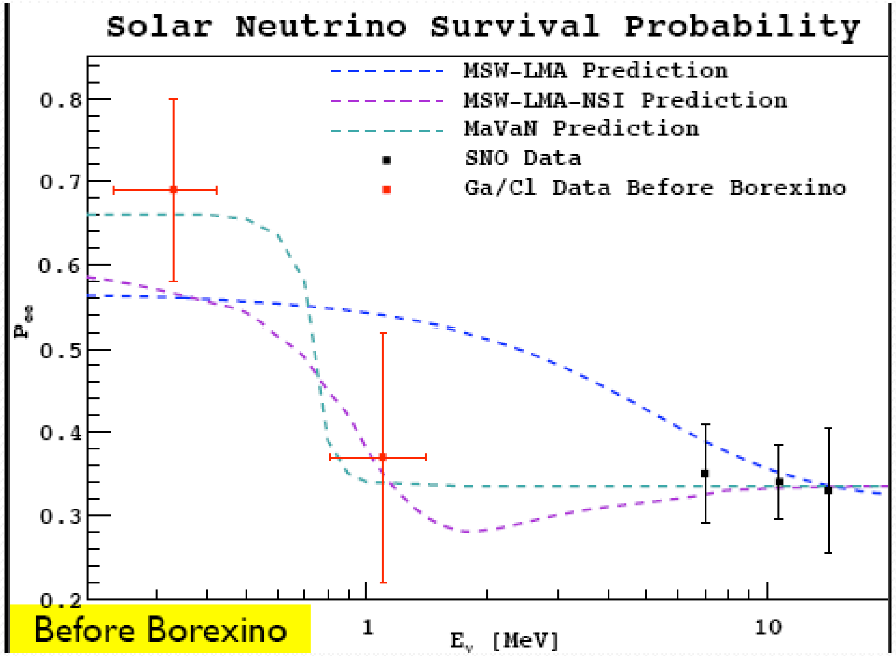
Systematic uncertainties

Source	Syst.error (1σ)
Tot. scint. mass	± 0.2%
Live Time	± 0.1%
Efficiency of Cuts	± 0.3%
Detector Resp.Function	± 6%
Fiducial Mass	± 6%
TOT	± 8.5%

$49 \pm 3_{stat} \pm 4_{sys}$ cpd/100 tons

	Expected rate (cpd/100 t)
No oscillation	75 ± 4
BPS07(GS98) HighZ	48 ± 4
BPS07(AGS05) LowZ	44 ± 4

No-oscillation hypothesis rejected at 4σ level



Under the assuptions of High-Z SSM (BPS 07) the ^{7}Be rate measurement corresponds to

P_{ee} (^{7}Be) = 0.56 0.1 (1σ)

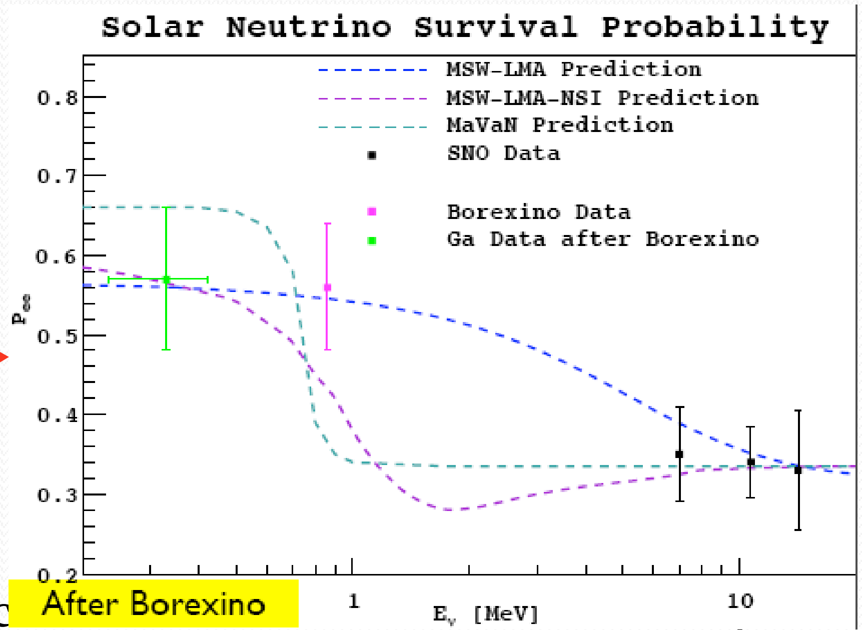
which is consistent with the number derived from the global fit to all solar and reactor experiments (S. Abe et al., arXiv: 0801.4589v2)

P_{ee} (^{7}Be) = 0.541 ± 0.017

We determine the survival probability for ^{7}Be and pp-ν_e, assuming BPS07 and **using input from all solar experiments** (Barger *et al.*, PR (2002) 88, 011302)

P_{ee} (^{7}Be) = 0.56 ± 0.08

P_{ee} (pp) = 0.57 ± 0.09



Neutrino magnetic moment

SM with $m_\nu = 0$: $\mu_\nu = 0$

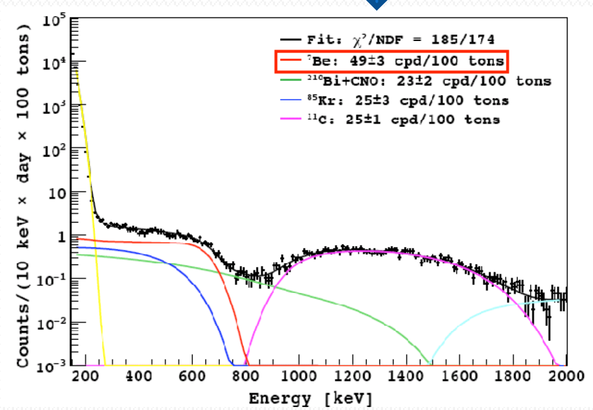
T – kinetic energy of scattered *e*
E_ν – neutrino energy

$$\left(\frac{d\sigma}{dT}\right)_W = \frac{2G_F^2 m_e}{\pi}\left[g_L^2 + g_R^2 + \left(1-\frac{T}{E_\nu}\right)^2 - g_L g_R \frac{m_e T}{E_\nu^2}\right]$$

SM with $m_\nu > 0$: $\mu_\nu > 0$, additional EM term influencing the cross section and thus the spectral shape

$$\left(\frac{d\sigma}{dT}\right)_{EM} = \mu_\nu^2 \frac{\pi\alpha_{em}^2}{m_e^2}\left(\frac{1}{T}-\frac{T}{E_\nu}\right)$$

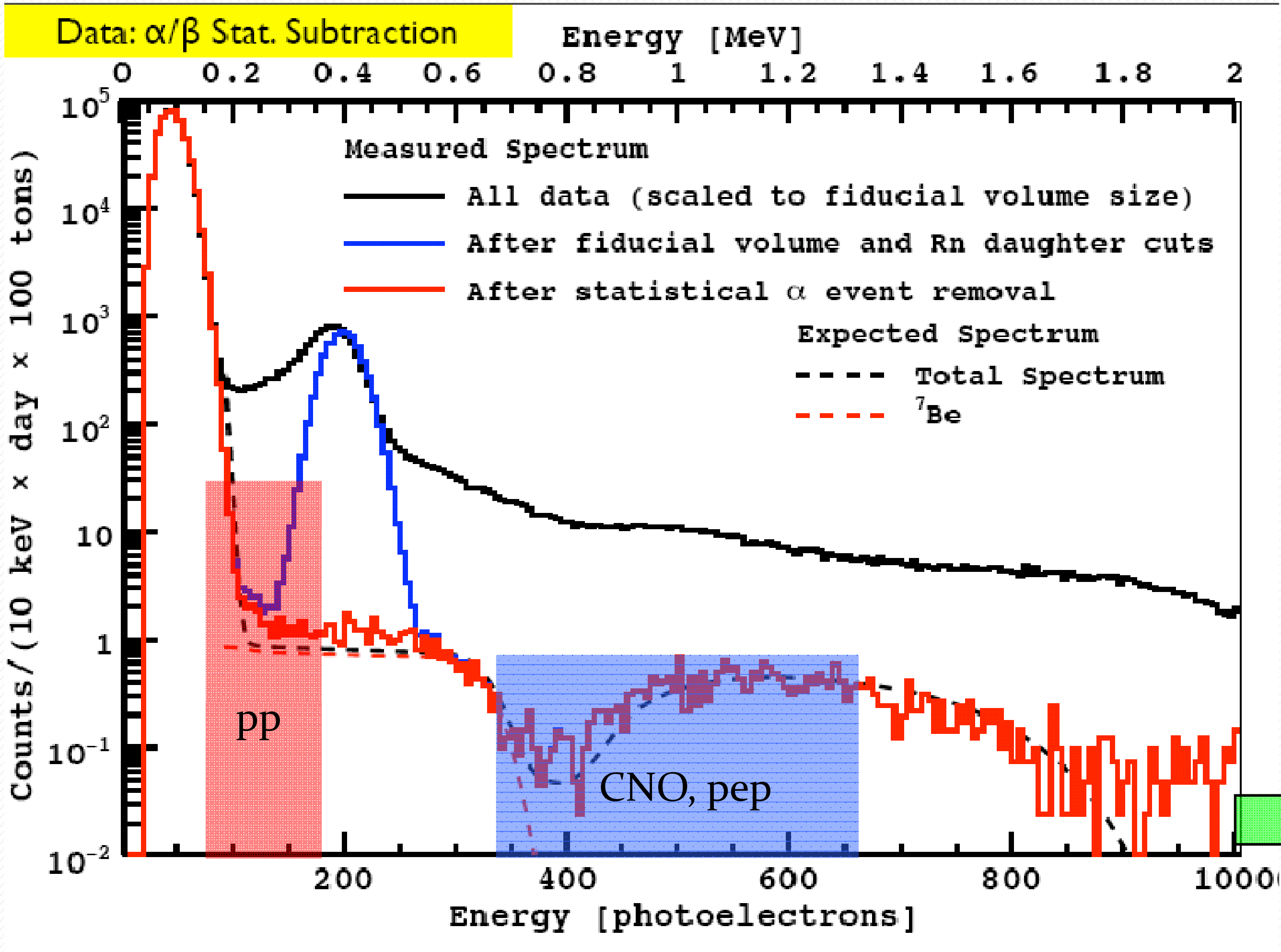
Sensitivity enhanced @ low energies

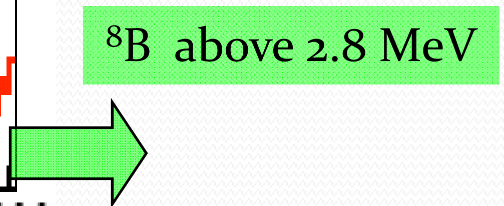


Estimate8	Method	90% C.L. 10^{-11} μ_B
SuperK	^{8}B above 5 MeV	< 11
Montanino *et al.*	^{7}Be (Borexino data)	< 8.4
GEMMA	Reactor anti-ν	< 5.8
Borexino	**^{7}Be**	**< 5.4**

Currently the best experimental limit!

What can Borexino say about other solar ν sources?





^{8}B-ν fluxes (arXiv 0808.2868)

- The first simultaneous measurement of solar–ν from the **transition region** (**^{7}Be-ν**) and from **the matter-enhanced oscillation region (^{8}B-ν);**
- The first measurement of **^{8}B-ν in real time below 5 MeV;**

Measurement of the solar ^{8}B neutrino flux with 246 live days of Borexino and observation of the MSW vacuum-matter transition

G. Bellini,[1] J. Benziger,[2] S. Bonetti,[1] M. Buizza Avanzini,[1] B. Caccianiga,[1] L. Cadonati,[3] F. Calaprice,[4] C. Carraro,[5] A. Chavarria,[4] F. Dalnoki-Veress,[4] D. D'Angelo,[1] H. de Kerret,[6] A. Derbin,[7] A. Etenko,[8] K. Fomenko,[9] D. Franco,[1] C. Galbiati,[4] S. Gazzana,[10] M. Giammarchi,[1] M. Goeger-Neff,[11] A. Goretti,[4] C. Grieb,[12,a] S. Hardy,[12] Aldo Ianni,[10] Andrea Ianni,[4] M. Joyce,[12] V. Kobychev,[13] G. Korga,[10] D. Kryn,[6] M. Laubenstein,[10] M. Leung,[4] T. Lewke,[11] E. Litvinovich,[8] B. Loer,[4] P. Lombardi,[1] L. Ludhova,[1] I. Machulin,[8] S. Manecki,[12] W. Maneschg,[14] G. Manuzio,[5] F. Masetti,[15] K. McCarty,[4] Q. Meindl,[11] E. Meroni,[1] L. Miramonti,[1] M. Misiaszek,[16,10] D. Montanari,[10,4] V. Muratova,[7] L. Oberauer,[11] M. Obolensky,[6] F. Ortica,[15] M. Pallavicini,[5] L. Papp,[10] L. Perasso,[1] S. Perasso,[5] A. Pocar,[4,b] R.S. Raghavan,[12] G. Ranucci,[1] A. Razeto,[10] P. Risso,[5] A. Romani,[15] D. Rountree,[12] A. Sabelnikov,[8] R. Saldanha,[4] C. Salvo,[5] S. Schönert,[14] H. Simgen,[14] M. Skorokhvatov,[8] O. Smirnov,[9] A. Sotnikov,[9] S. Sukhotin,[8] Y. Suvorov,[1,8] R. Tartaglia,[10] G. Testera,[5] D. Vignaud,[6] R.B. Vogelaar,[12] F. von Feilitzsch,[11] M. Wojcik,[16] M. Wurm,[11] O. Zaimidoroga,[9] S. Zavatarelli,[5] and G. Zuzel[14]

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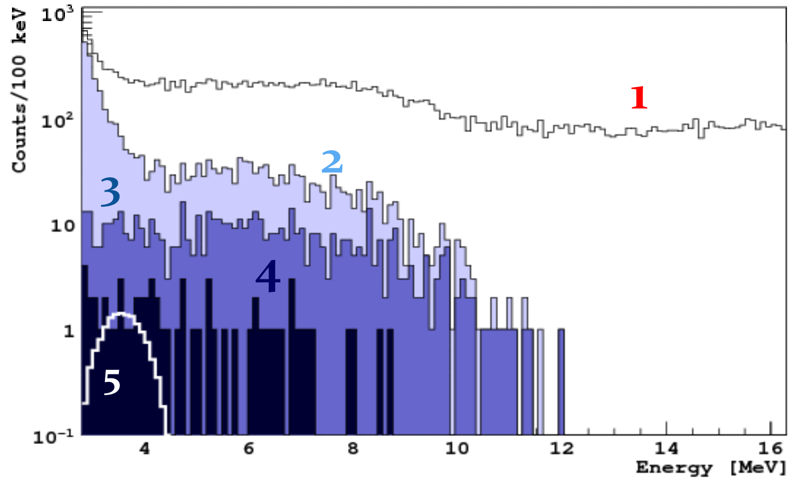
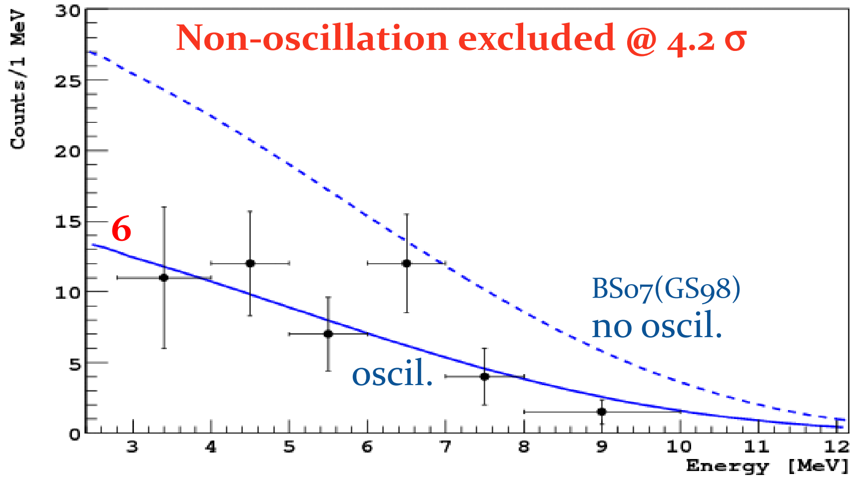
(Dated: August 21, 2008)

submitted to PRL

^{8}B-ν fluxes (see arXiv 0808.2868 for details)

Cut	Counts 2.8-16.3 MeV	Counts 5.0-16.3 MeV	Indx
None	20449	14304	1
Muon cut	3363	1135	2
Neutron cut	3280	1114	
FV cut	567	372	3
Cosmogenic cut	71	26	
^{10}C removal	65	26	
^{214}Bi removal	62	26	4
Expected ^{208}Tl	14 ± 3	0	5
Measured ^{8}B-ν	48 ± 8	26 ± 5	6
BS07(GS98) ^{8}B-ν	50 ± 5	25 ± 3	
BS07(AGS05) ^{8}B-ν	40 ± 4	20 ± 2	





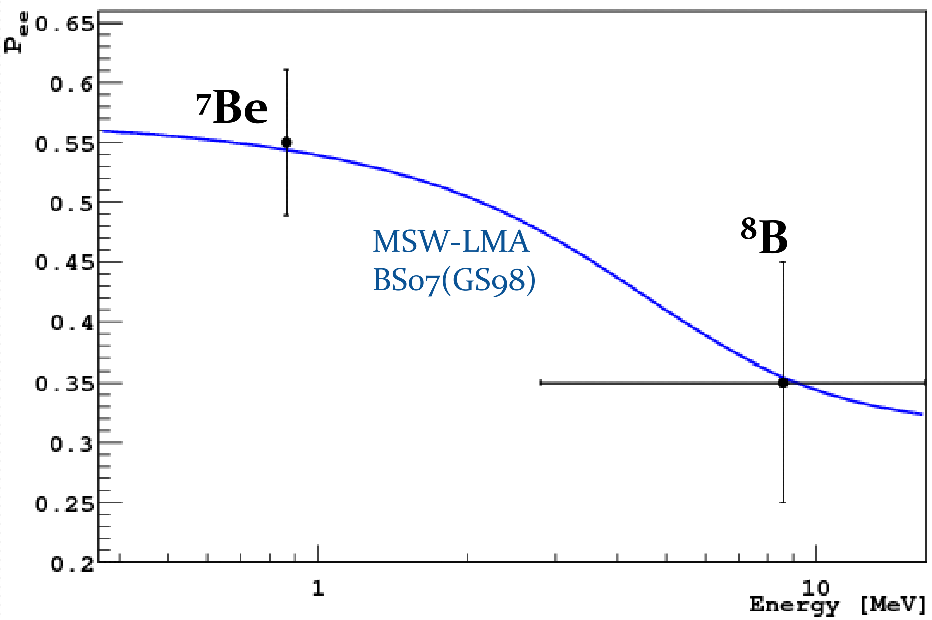
^{8}B-ν fluxes (see arXiv 0808.2868 for details)

First real-time measurement above 2.8 MeV:

$$\Phi_{>2.8\,MeV} = (0.26 \pm 0.04_{\text{stat}} \pm 0.02_{\text{sys}})\ \text{counts/day /100 tons}$$

Above 5 MeV in agreement with SNO and SuperK:

$$\Phi_{>2.8\,MeV} = (0.14 \pm 0.03_{\text{stat}} \pm 0.01_{\text{sys}})\ \text{counts/day /100 tons}$$



BS05 (GS98), HighZ

First simultaneous measurement in both vacuum-dominated and matter-enhanced regions

Conclusions

DONE

- Borexino performed the first real-time measurement **of solar-ν below the barrier of natural radioactivity** (4 MeV);
- The two measurements reported for **^{7}Be-ν** favor MSW-LMA solution;
- The first real-time measurement of **^{8}B-ν below 5 MeV (above 2.8 MeV)** ;
- The first **simultaneous** measurement of solar neutrinos from the **transition region** (^{7}Be-ν) and from **the matter-enhanced oscillation region** (^{8}B-ν);
- **Best limits for pp- and CNO-ν,** combining information from all solar and reactor experiments;

TO BE DONE

- Precision measurement (at or below the level of 5%) of the ^{7}Be-ν rate;
- Check the 7% seasonal variation of the neutrino flux (confirm solar origin);
- Under study: measurement of the CNO, pep and high-energy pp neutrinos;
- Strong potential in antineutrinos (geoneutrinos, reactor, from the Sun) and in supernovae neutrinos and antineutrinos;

In memoriam:

- Engin Arik
- Berkol Dogan
- Engin Abat
- Senel Boydag
- Iskender Hikmet
- Mustafa Fidan

Thank you!!

Additional slides

Constraints on *pp*- & CNO-ν fluxes after the ^{7}Be measurement

- It is possible to combine the results obtained by Borexino on ^{7}Be flux with those obtained by other experiments to constraint the fluxes of pp and CNO ν_e;
- The measured rate in Clorine and Gallium experiments can be written as:

$$R_k = \sum_{i,k} f_i R_{i,k} P_{ee}^{i,k}$$

$$k = \text{Homestake, Gallex}$$

$$i = pp, pep, CNO, {}^7Be, {}^8B$$

$$f_i = \frac{\phi_i(\text{measured})}{\phi_i(\text{predicted})}$$

$$R_{i,k} = \text{expected rate of source "i" in experiment "k" (no oscill.)}$$

$$P_{ee}^{e,k} = \text{average survival probability for source "i" in experiment "k"}$$

- $R_{i,k}$ and $P_{ee}^{i,k}$ are calculated in the hypothesis of **high-Z SSM and MSW LMA**, ;
- R_k are the rates actually measured by Clorine and Gallium experiments;
- **f_{8B} = 0.87** ± **0.07**, measured by **SNO** and **SuperK;**
- **f_{7Be} = 1.02 ±0.10** is given by **Borexino results**;
- Performing a χ^2 based analysis with the additional luminosity constraint;

$$f_{pp} = 1.005^{+0.008}_{-0.020} \quad (1\sigma)$$

$$f_{CNO} < 3.80 \quad (90\% \ C.L.)$$

Which is the best determination of pp flux (with luminosity constraint)

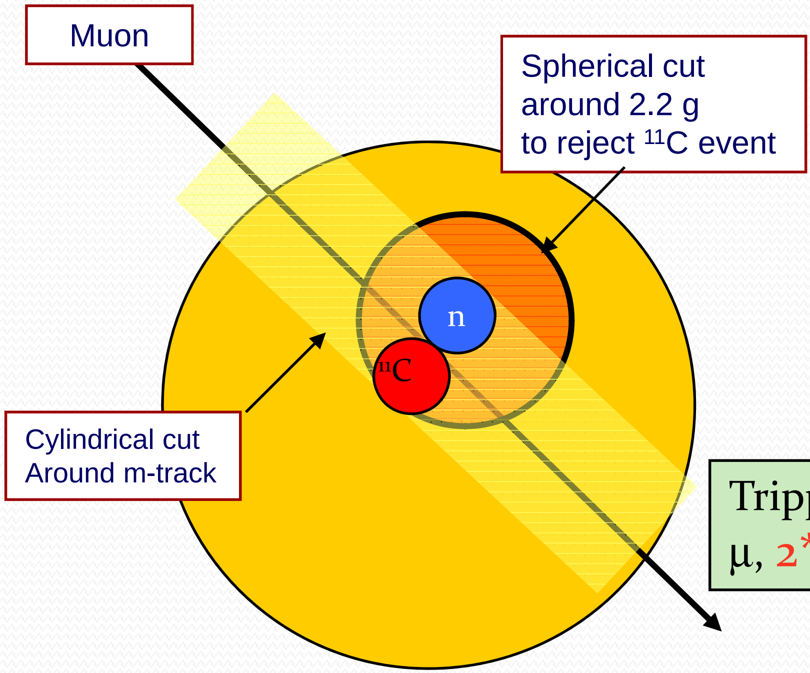
Future: pep and CNO-ν fluxes

The main background for pep and CNO analysis is ^{11}C

$$\mu + {}^{12}C \rightarrow \mu + {}^{11}C + n \ (90\%)$$

τ ~ 30 min!! τ ~ 260 μs

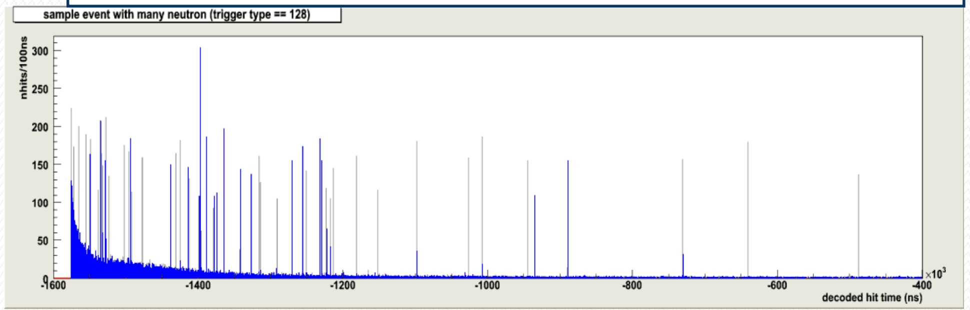
$$^{11}C \rightarrow {}^{11}B + e^+ + \nu_e \qquad n + p \rightarrow d + \gamma \ (2.2\ \text{MeV})$$



Tripple coincidence:
μ, 2*511 keV, γ (2.2 MeV)

- electronics improvement to detect all the neutrons produced by a muon
  - Changes in the electronics (Dec 07): after each muon, **1.6 ms gate opened**
  - FADC implementation in parallel;

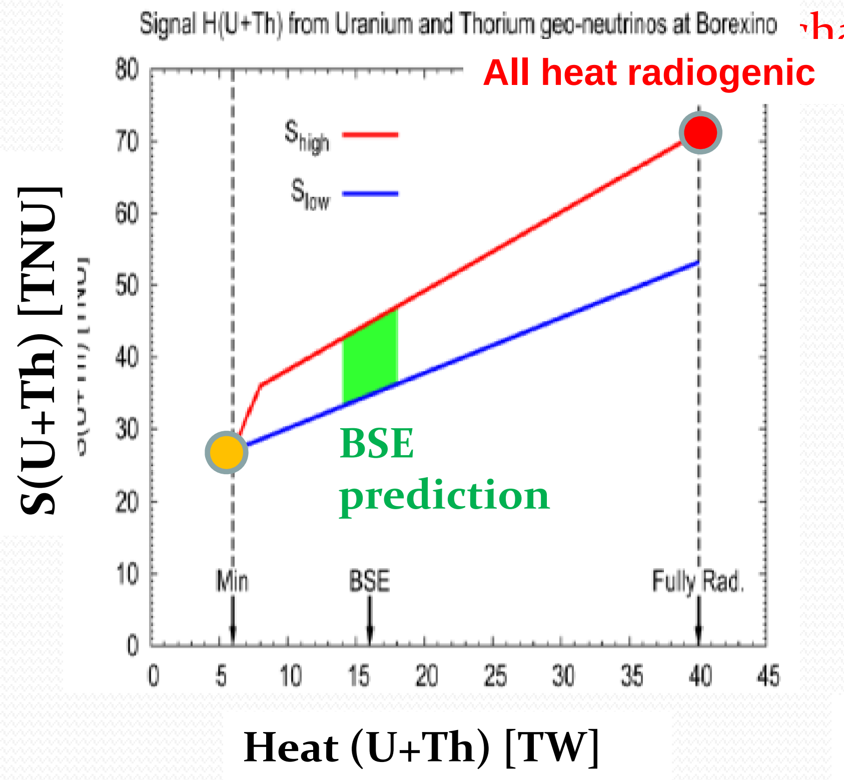
Example with several tens of neutrons detected:



Now: 87% ¹¹C rejection

Still in progress

Borexino potential on geo-neutrinos

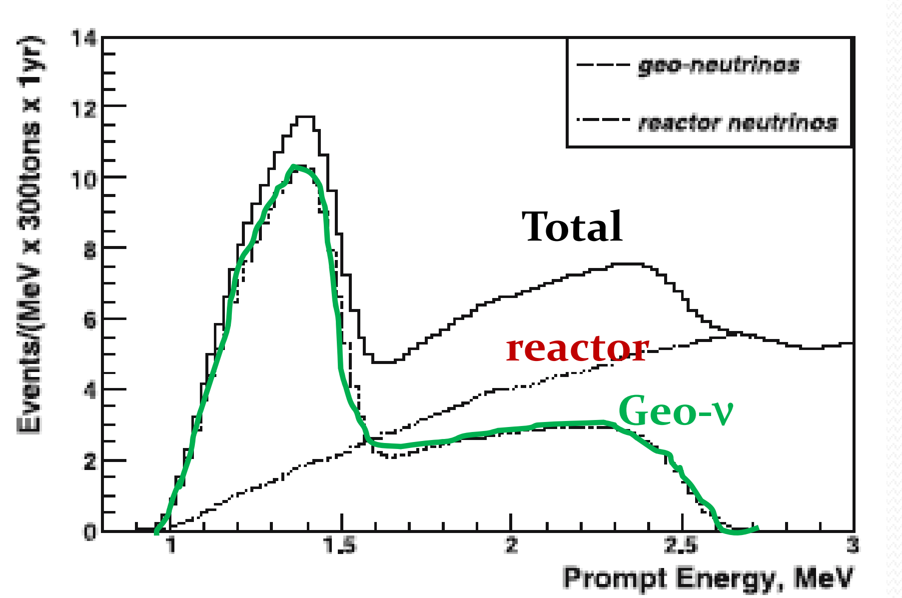


Mantovani et al., TAUP 2007

TNU = 1 event / 10^{32} target proton / year

Np (Borex) = 1.8 10^{31} target proton

Prompt signal energy spectrum (model)



5.7 events from reactors (in geo-ν E range)

BSE: 6.3 events from geoneutrinos

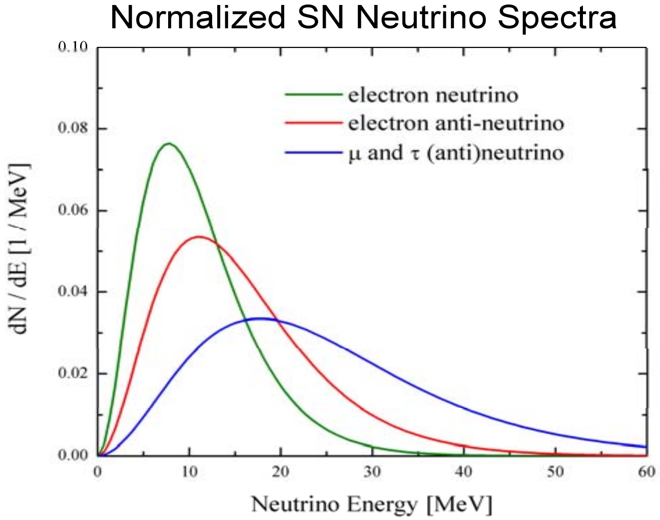
(per year and 300 tons, ε = 80%, 1-2.6 MeV)

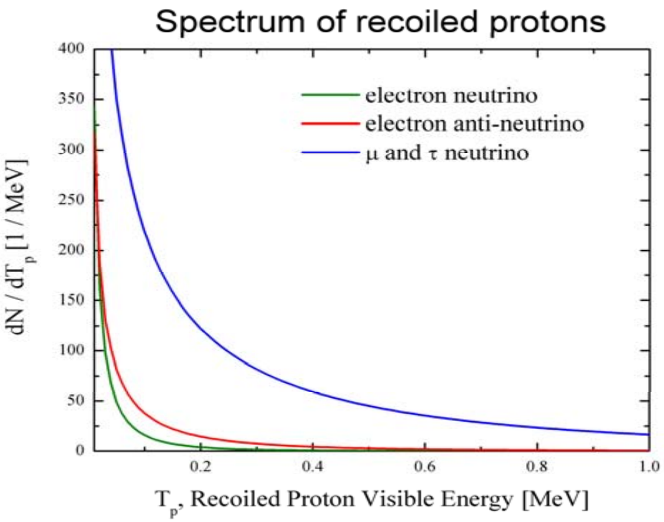
(Balata *et al.,* 2006, ref. model Mantovani *et al.*, 2004)

3σ evidence of geoneutrinos expected in 2 years of data

Borexino potential on supernovae neutrinos

Standard SN @ 10kpc





Borexino E_{tresh} = 0.25 MeV
target mass 300 t

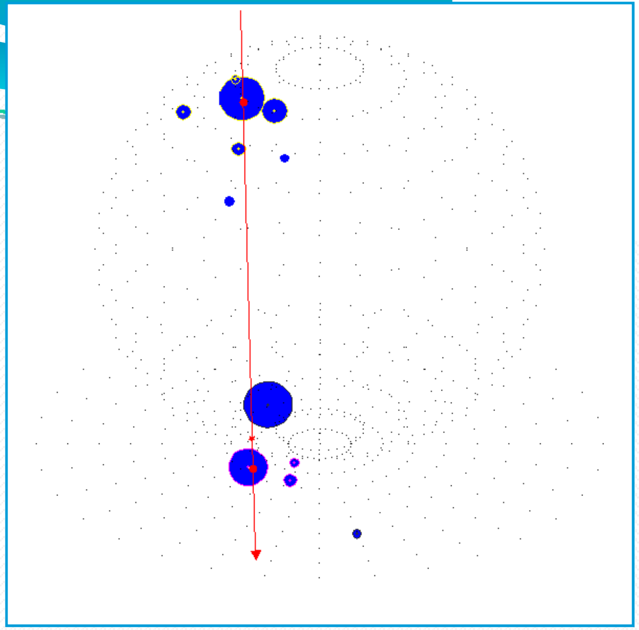
Detection channel	N events
ES (E_ν > 0.25 MeV)	5
Electron anti-neutrinos (E_ν > 1.8 MeV)	78
ν-p ES (E_ν > 0.25 MeV)	52
$^{12}C(\nu,\nu)^{12}C^*$ (Eγ = 15.1 MeV)	18
$^{12}C(\text{anti-}\nu,e^+)^{12}B$ ($E_{anti-\nu}$ > 14.3 MeV)	3
$^{12}C(\nu,e^-)^{12}N$ (E_ν > 17.3 MeV)	9

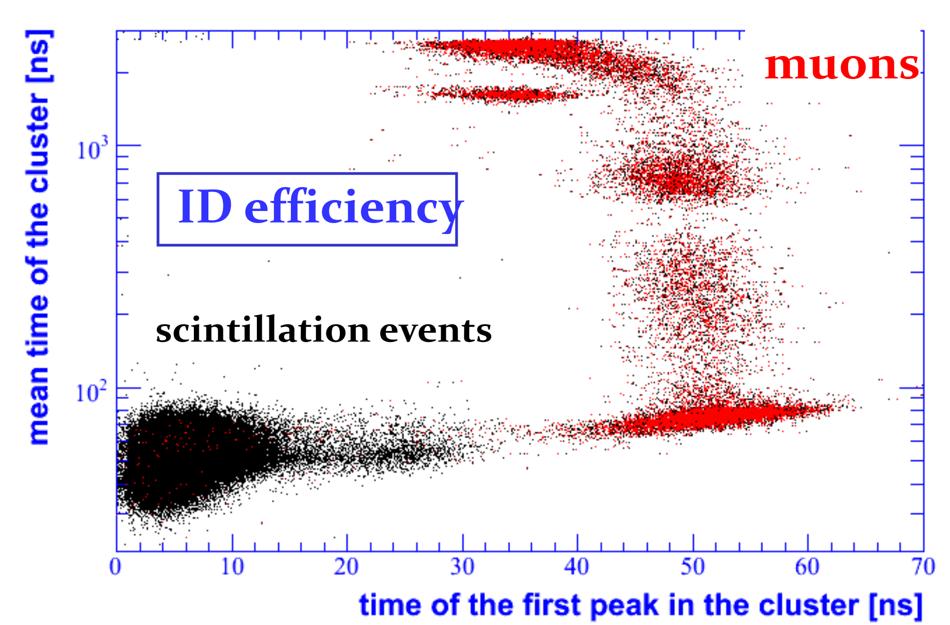
Can be used as an early alarm

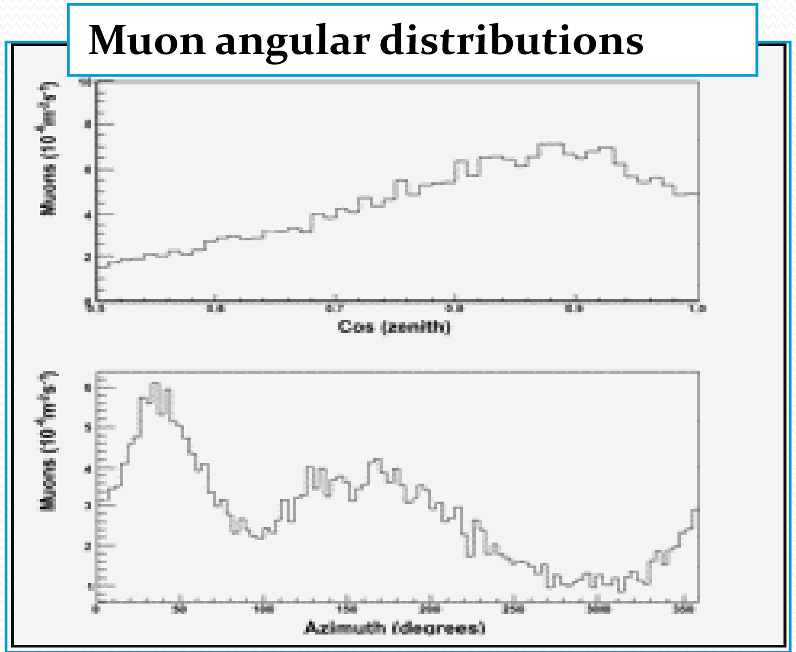
Borexino will enter SNEWS soon

Cosmic μ's

- μ are identified by the OD and by the ID
  - OD eff: ~ 99%
  - ID analysis based on pulse shape variables
    - Deutsch variable: ratio between light in the concentrator and total light
    - Cluster mean time, peak position in time
  - Estimated overall rejection factor > 10^4 (still preliminary)
- **After cuts, m not a relevant background for ^{7}Be**
  - **Residual background: < 1 count /day/ 1 00 t**



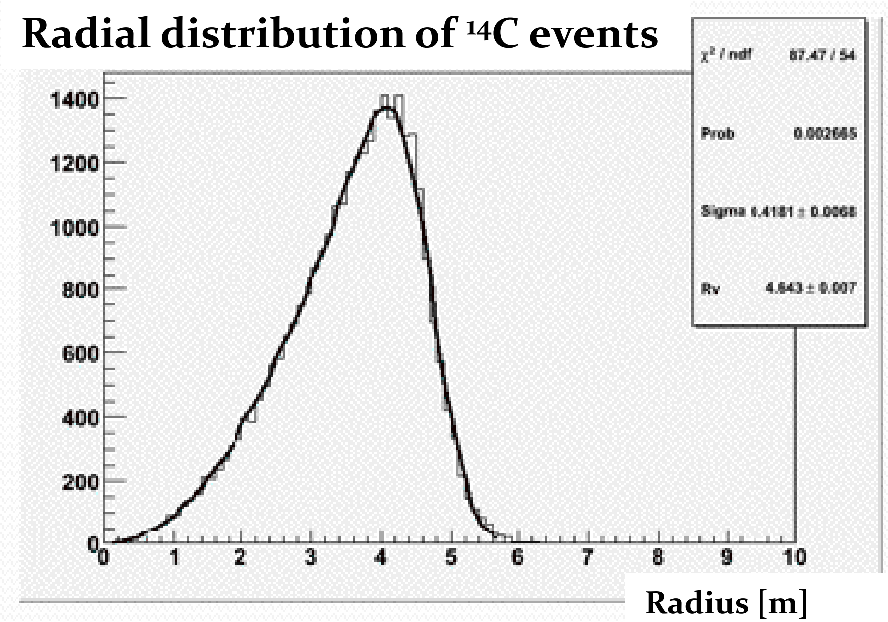




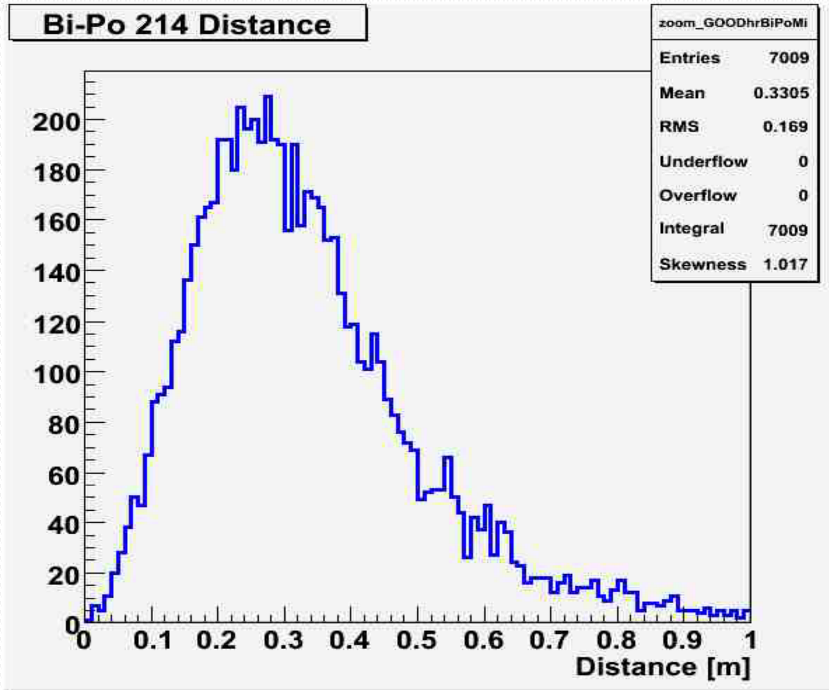
Muon flux:(1.21±0.05)$h^{-1}m^{-2}$

Position reconstruction algorithms

- Base on time of flight fit to hit-time distribution
- developed with MC, tested and validated in Borexino prototype CTF
- cross checked and tuned in Borexino on selected events (^{14}C, ^{214}Bi-^{214}Po, ^{11}C)

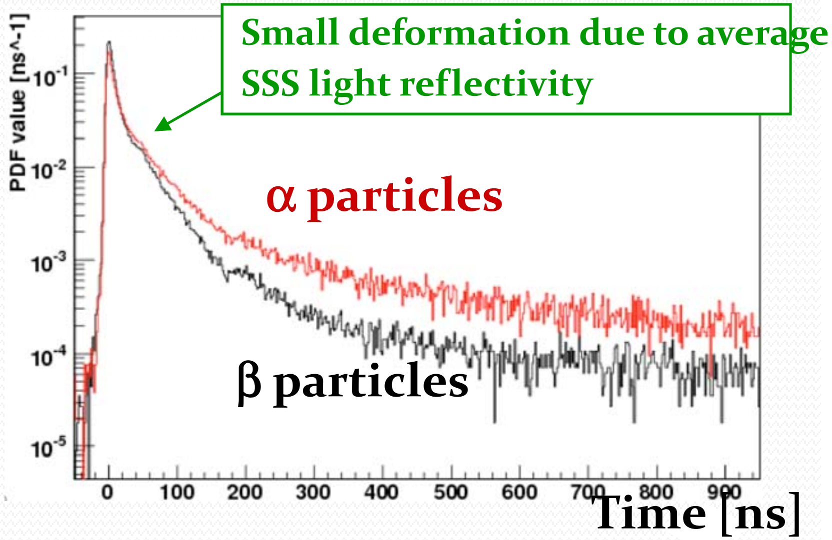


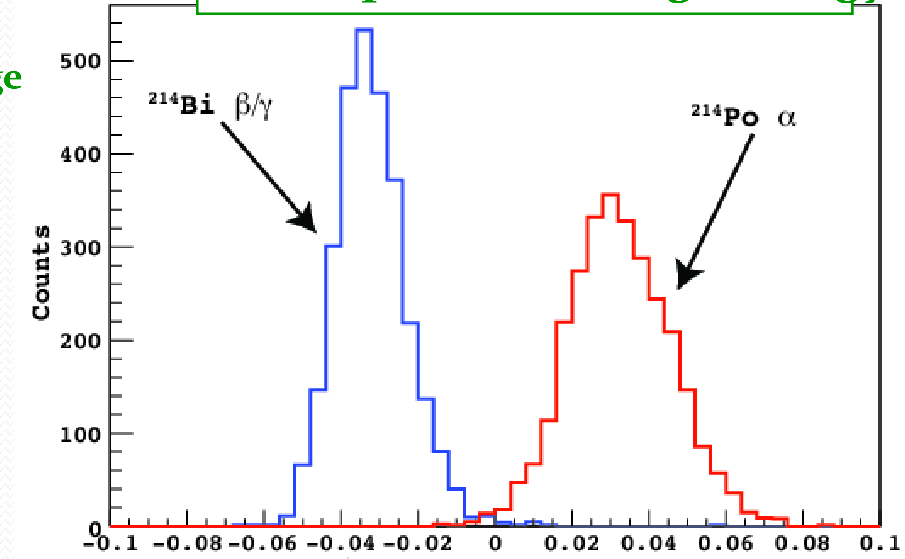
^{14}C "bound" in the scintillator: homogeneous
The fit is compatible with the expected
r^2-like shape with R=4.25m.



Spatial resolution:
35 cm at 200 keV
16 cm at 500 keV
(scaling as $N_{p.e.}^{-1/2}$)

α/β discrimination





$\overline{\alpha}_i, \overline{\beta}_i$ –> average pulse shapes

$$P_i = \frac{(\overline{\alpha}_i - \overline{\beta}_i)}{(\overline{\alpha}_i + \overline{\beta}_i)}$$ –> for i-time interval of 2 ns

$$G = \sum_i P_i S_i$$

S_i –> signal shape within a given Δt (2 ns)

Techinques towards low-radioactivity

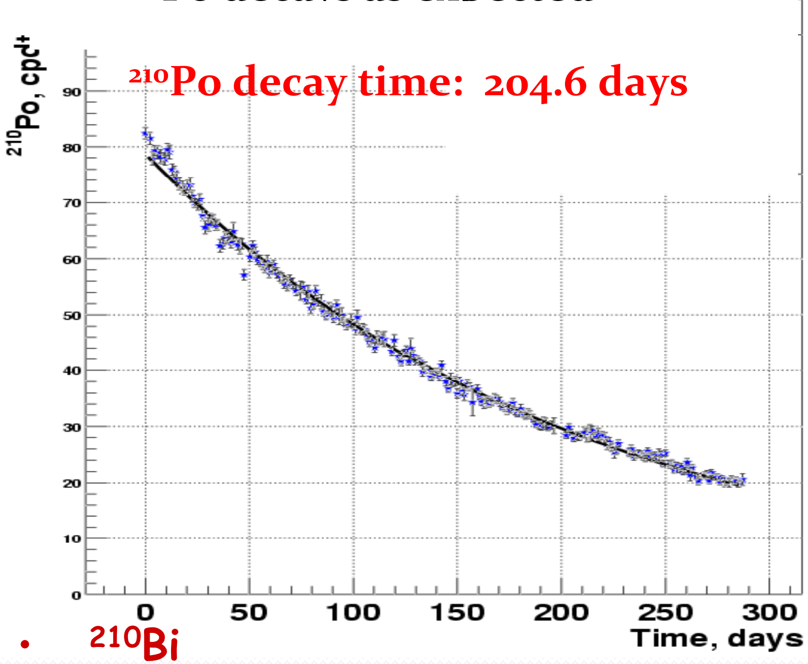
- Low background nylon vessel fabricated in hermetically sealed low radon clean room (~1 yr)
- Rapid transport of scintillator solvent (PC) from production plant to underground lab to avoid cosmogenic production of radioactivity (^{7}Be)
- Underground purification plant to distill scintillator components.
- Gas stripping of scintlllator with special nitrogen free of radioactive ^{85}Kr and ^{39}Ar from air
- All materials electropolished SS or teflon, precision cleaned with a dedicated cleaning module

Background: ^{210}Po

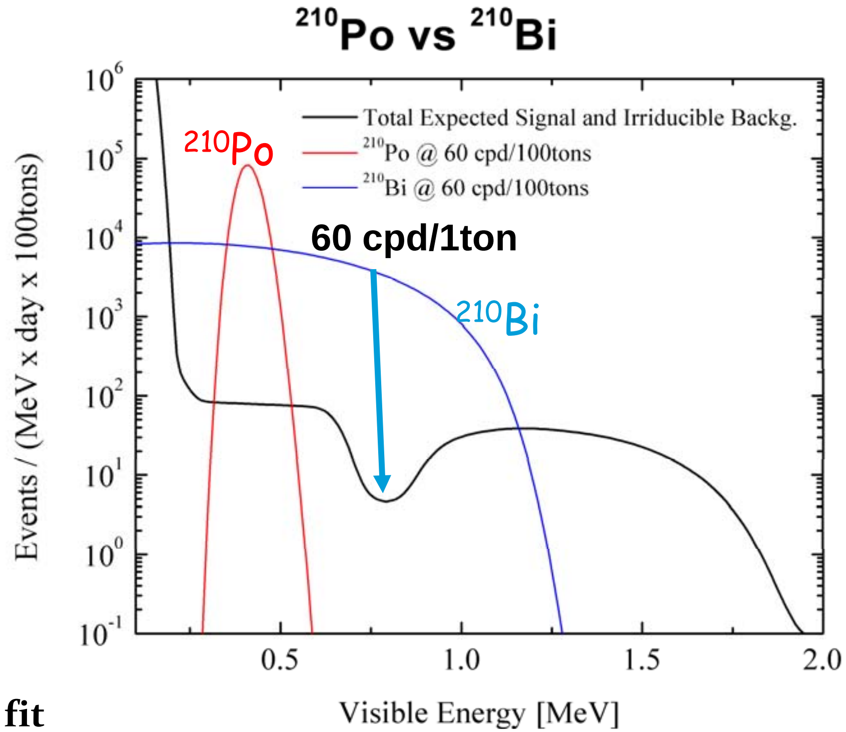
End of ^{238}U chain :

	^{210}Pb	->	^{210}Bi	->	^{210}Po	->	^{206}Pb
		β^-(61 keV)		β^-(1.2MeV)		α	
$t_{1/2}$	22.3 y		5.01 d		138.38 d		stable

- **Not in equilibrium with ^{210}Pb !**
- **^{210}Po decays as expected**



- The bulk ^{238}U and ^{232}Th contamination is negligible
- The ^{210}Po background is NOT related neither to ^{238}U nor to ^{210}Pb contamination

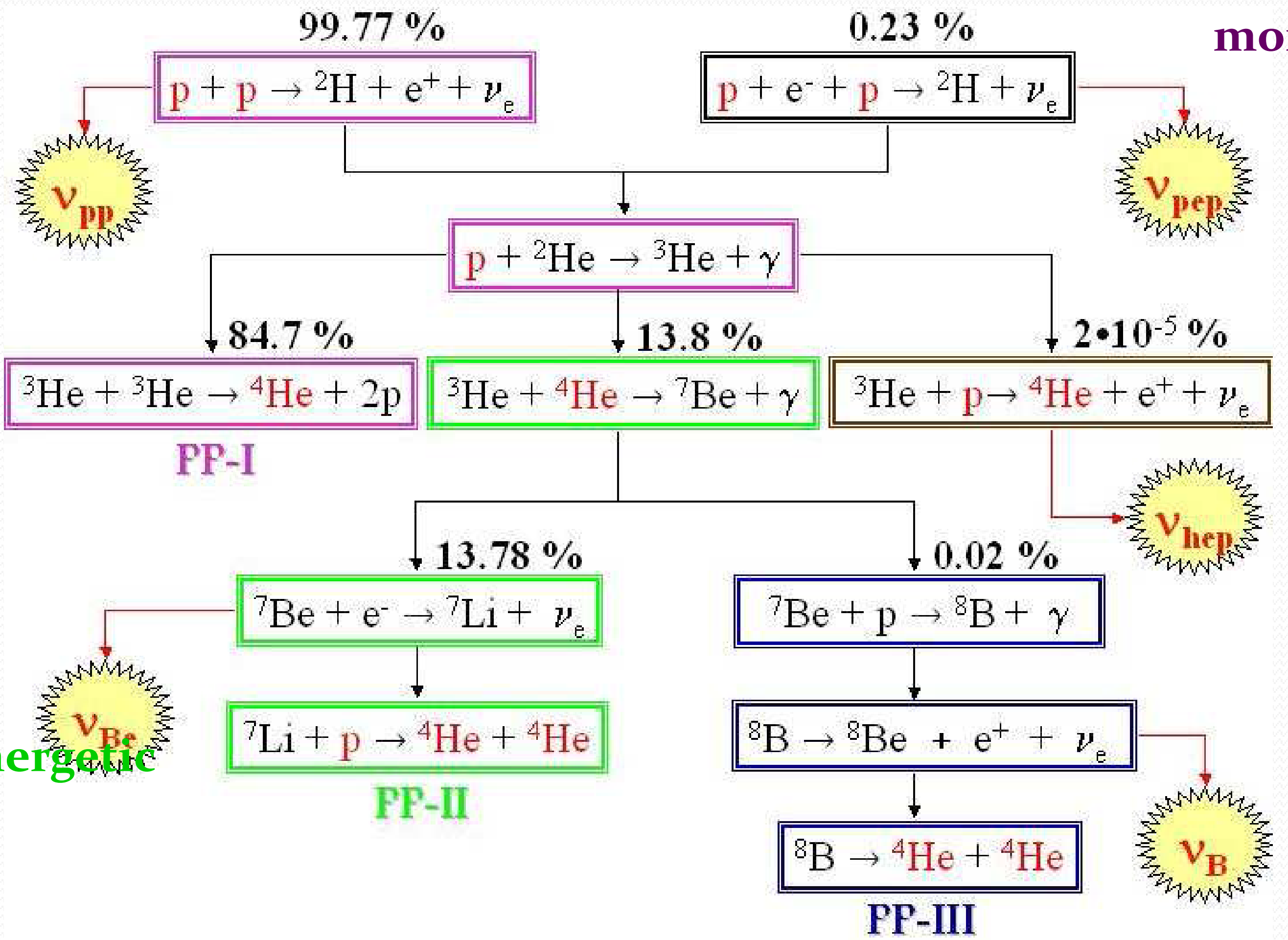


- ^{210}Bi

no direct evidence----> free parameter in the total fit

cannot be disentangled, in the ^{7}Be energy range, from the CNO

Proton-proton cycle: the main energy source in the Sun



monoenergetic

monoenergetic