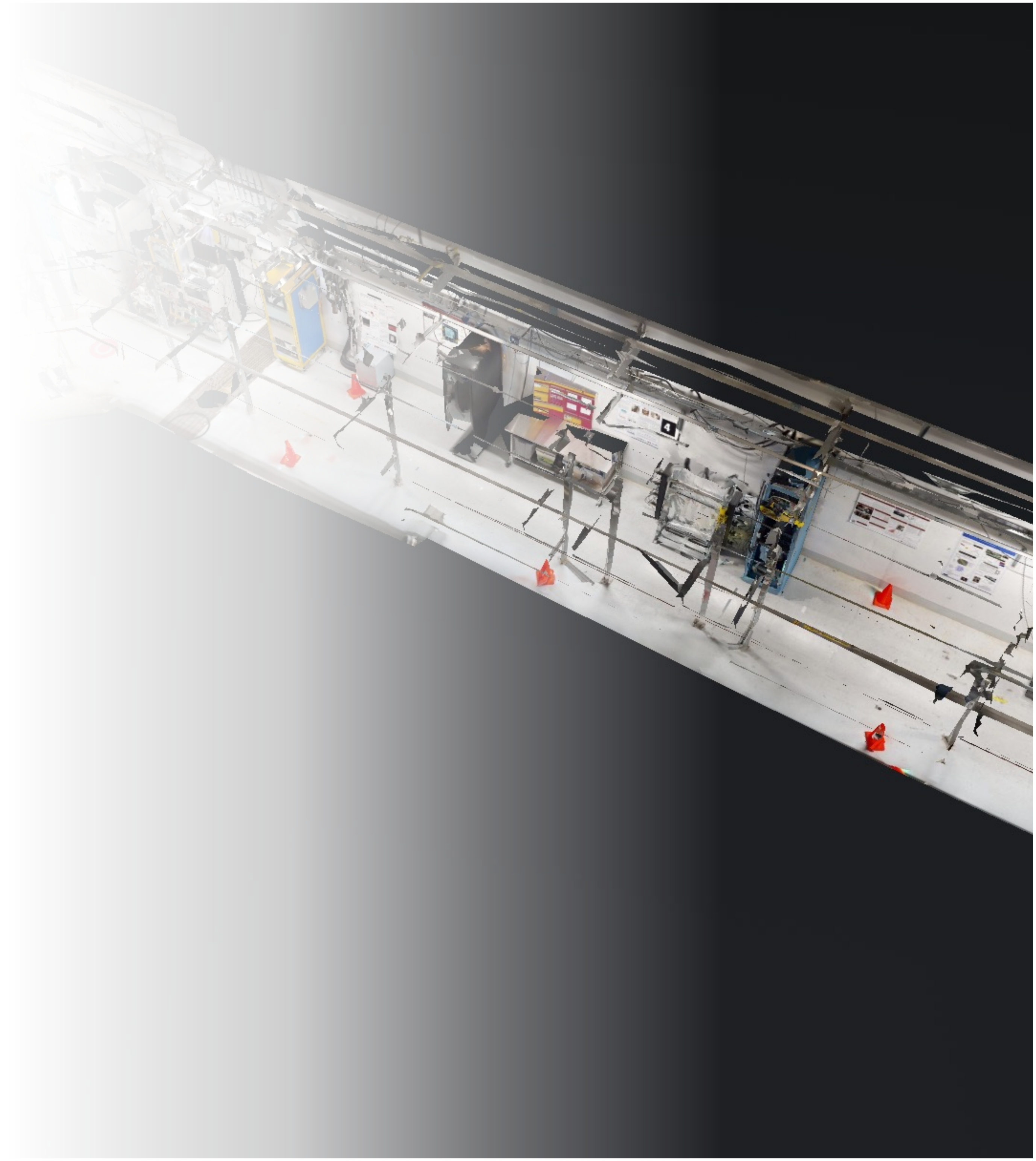
COHERENT ELASTIC NEUTRINO-NUCLEUS SCATTERING



Kate Scholberg,
Duke University

INSS
August 2021

OUTLINE

- Neutrinos and neutrino interactions
- Coherent elastic neutrino-nucleus scattering (CEvNS)
- Why measure it? Physics motivations
  (short and long term)
- How to measure CEvNS
- The COHERENT experiment at the SNS
  - First measurements
  - Status and prospects
- Worldwide status of CEvNS experiments

Neutrino interactions over a range of energies

Deep Inelastic Scattering

keV MeV GeV TeV PeV

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Interactions with nucleons inside nuclei, often disruptive, hadroproduction

Neutrino interactions over a range of energies

Interactions with nuclei and electrons, minimally disruptive of the nucleus

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keV MeV GeV TeV PeV

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Coherent elastic neutrino-nucleus scattering

Interactions with nuclei and electrons, minimally disruptive of the nucleus

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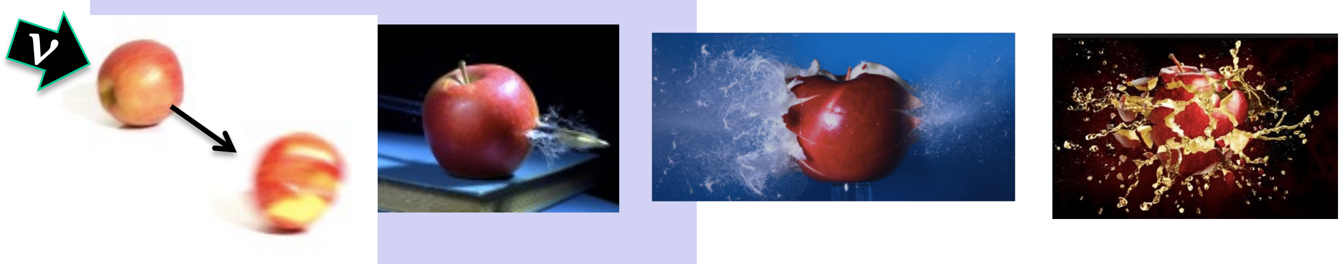
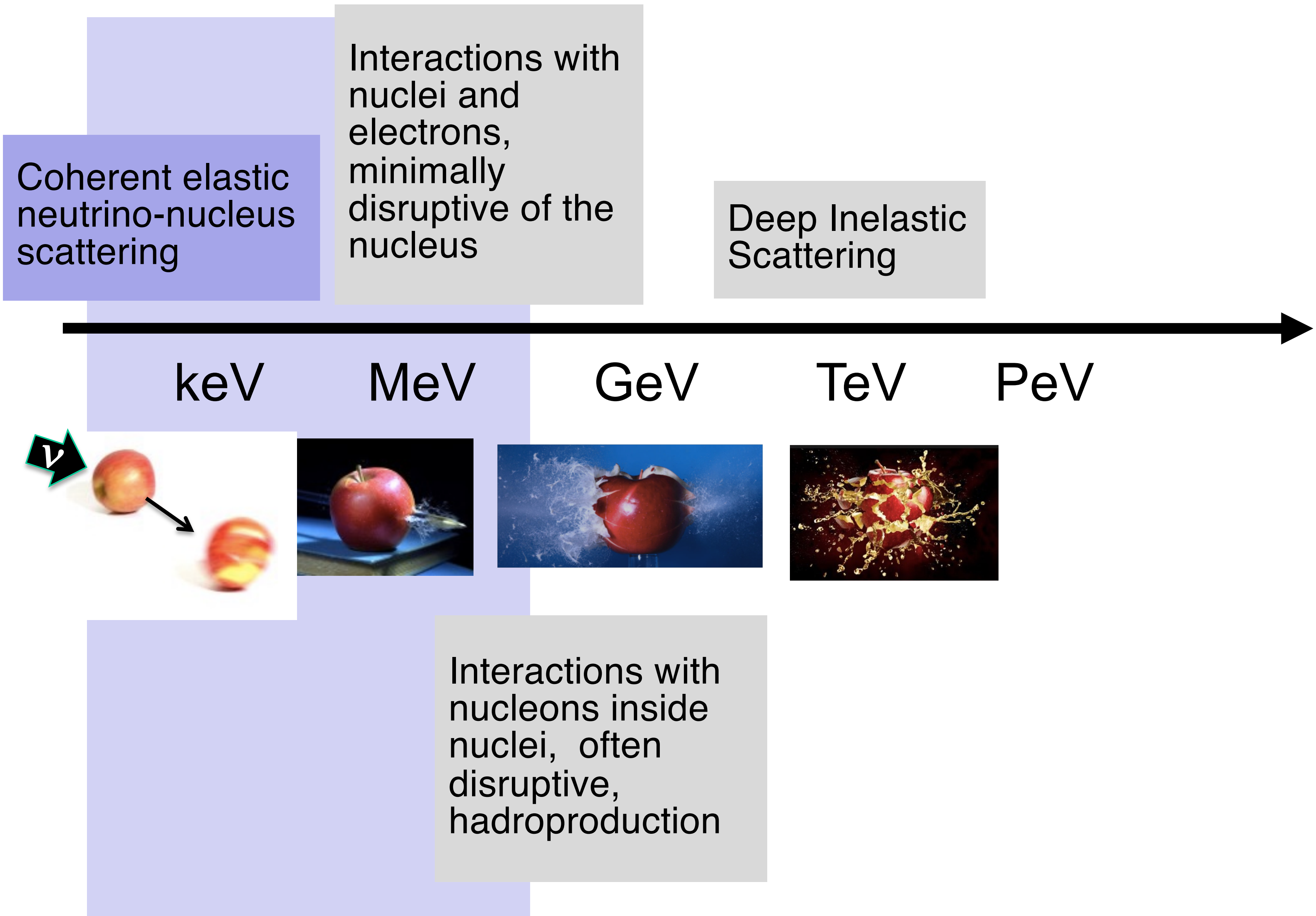


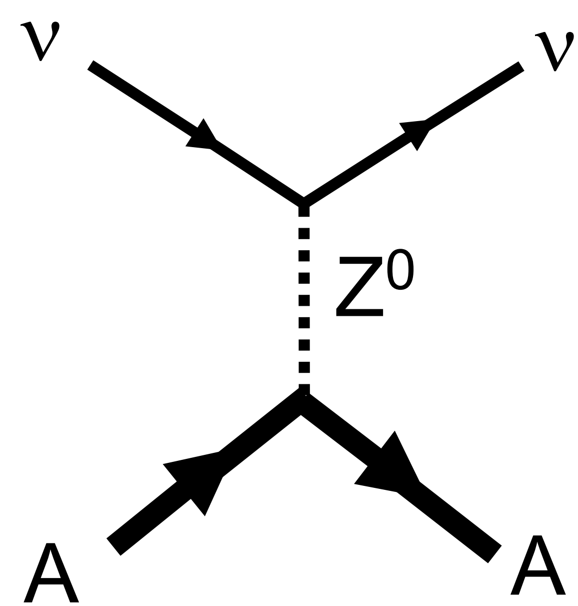
Interactions with nucleons inside nuclei, often disruptive, hadroproduction

This is the ***gentlest*** interaction of a neutrino with a nucleus

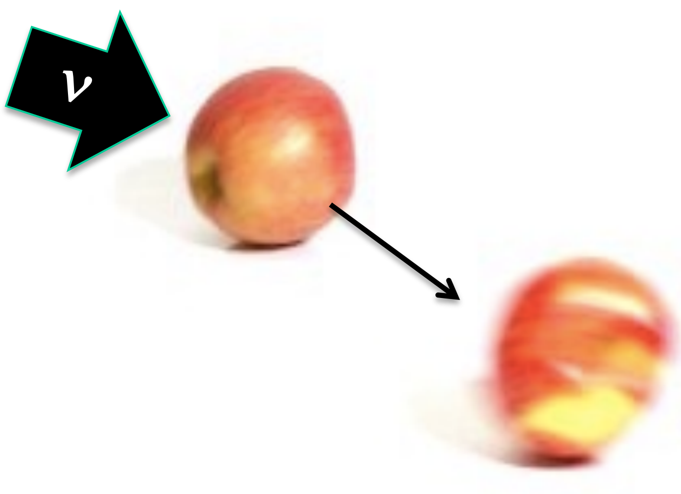


Coherent elastic neutrino-nucleus scattering (CEvNS)

$$\nu + A \rightarrow \nu + A$$



A neutrino smacks a nucleus via exchange of a Z, and the nucleus recoils as a whole; **coherent** up to $E_\nu \sim 50$ MeV



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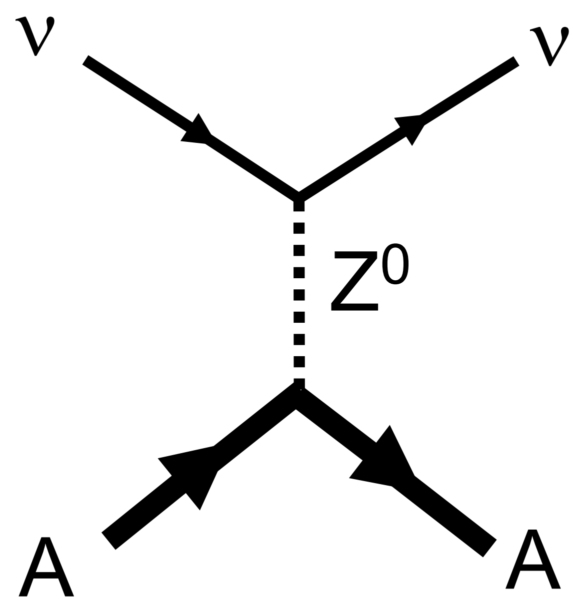

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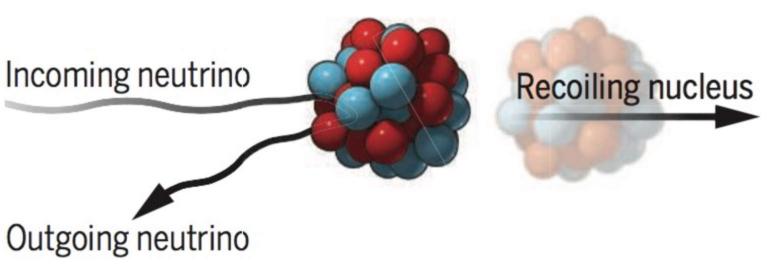


Nucleon wavefunctions in the target nucleus are **in phase with each other** at low momentum transfer

For $QR << 1$, [total xscn] ~ A^2 * [single constituent xscn]

A: no. of constituents

Image: J. Link *Science* Perspectives

First proposed 44 years ago!

PHYSICAL REVIEW D VOLUME 9, NUMBER 5 1 MARCH 1974

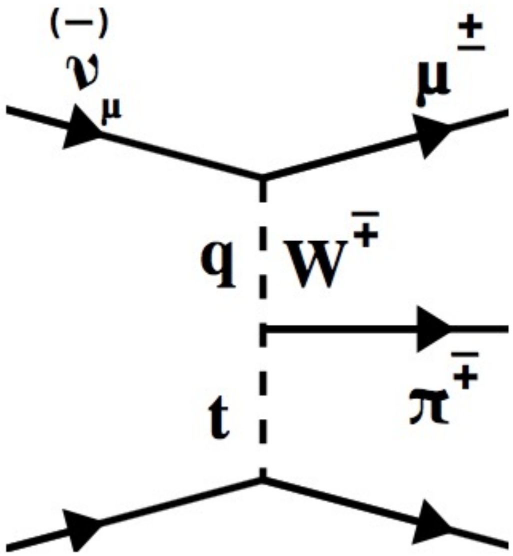
Coherent effects of a weak neutral current

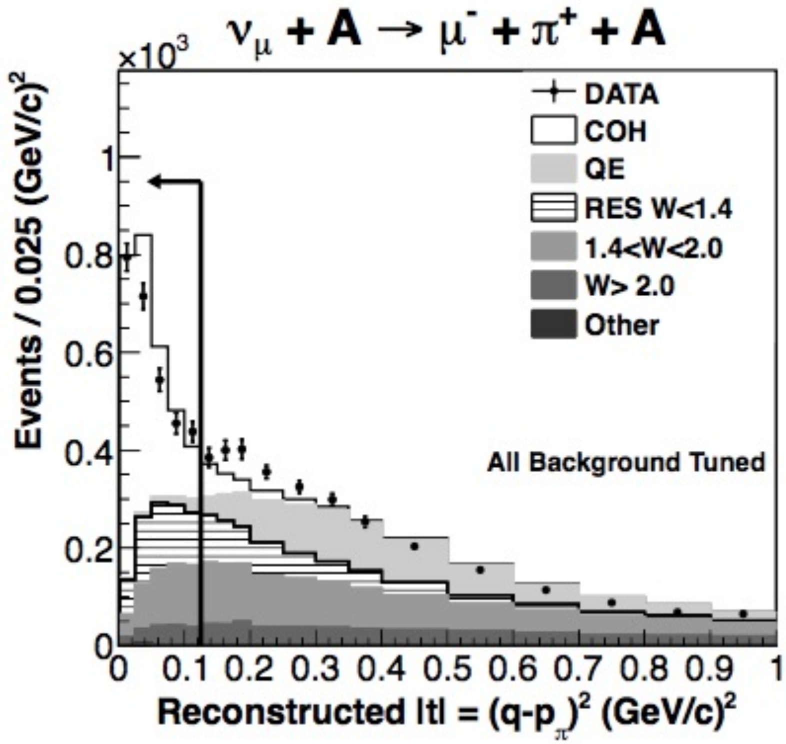
Daniel Z. Freedman†
National Accelerator Laboratory, Batavia, Illinois 60510
and Institute for Theoretical Physics, State University of New York, Stony Brook, New York 11790
(Received 15 October 1973; revised manuscript received 19 November 1973)

Our suggestion may be an act of hubris, because the inevitable constraints of interaction rate, resolution, and background pose grave experimental difficulties for elastic neutrino-nucleus scattering. We will discuss these problems at the end of this note, but first we wish to present the theoretical ideas relevant to the experiments.

Also: D. Z. Freedman et al., "The Weak Neutral Current and Its Effect in Stellar Collapse", Ann. Rev. Nucl. Sci. 1977. 27:167-207

This is ***not*** coherent pion production, a strong interaction process ***(inelastic)***





A. Higuera et. al, MINERvA collaboration, PRL 2014 113 (26) 2477

How do you pronounce "CEvNS"?

A. "KEVENS"

B. "KENZ"

C. "KENSE"

D. "SEVENS"

E. "SENSE"

F. "SENZ"

\begin{aside}

Literature has CNS, CNNS, CENNS, ...

- I prefer including "E" for "elastic"... otherwise it gets frequently confused with coherent pion production at ~GeV neutrino energies
- I'm told "NN" means "nucleon-nucleon" to nuclear types
- CEνNS is a possibility but those internal Greek letters are annoying

➔**CEvNS**, pronounced "sevens"...

spread the meme!

\end{aside}

Standard Model prediction for CEvNS differential cross section

E_ν: neutrino energy
T: nuclear recoil energy
M: nuclear mass
$Q = \sqrt{(2\,M\,T)}$:
momentum transfer

(probability of kicking a nucleus with recoil energy T)

Fermi constant (SM parameter)

$$\frac{d\sigma}{dT} \simeq \frac{\boxed{G_F^2} M}{2\pi} \frac{Q_W^2}{4} F^2(Q) \left(2 - \frac{MT}{E_\nu^2}\right)$$

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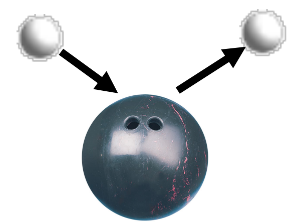
kinematics:
ping-pong ball hits bowling ball



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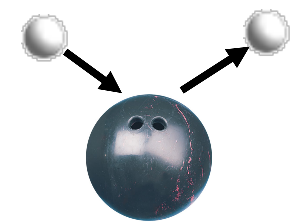
Form factor: F=1 ➔ full coherence



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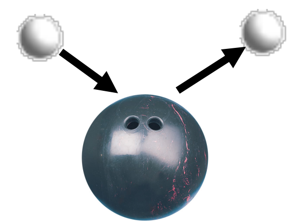

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weak nuclear charge

$$Q_W = N - (1 - 4\sin^2\theta_W)Z$$

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weak nuclear charge

No. of neutrons

No. of protons

$$Q_W = N - \boxed{(1 - 4\sin^2\theta_W)}Z$$

$\sin^2\theta_W = 0.231$, so protons unimportant $\Longrightarrow$ $Q_W \propto N$

$$\frac{d\sigma}{dT} \simeq \frac{G_F^2 M}{2\pi} \frac{\boxed{Q_W^2}}{4} \boxed{F^2(Q)} \left(2 - \frac{MT}{E_\nu^2}\right)$$

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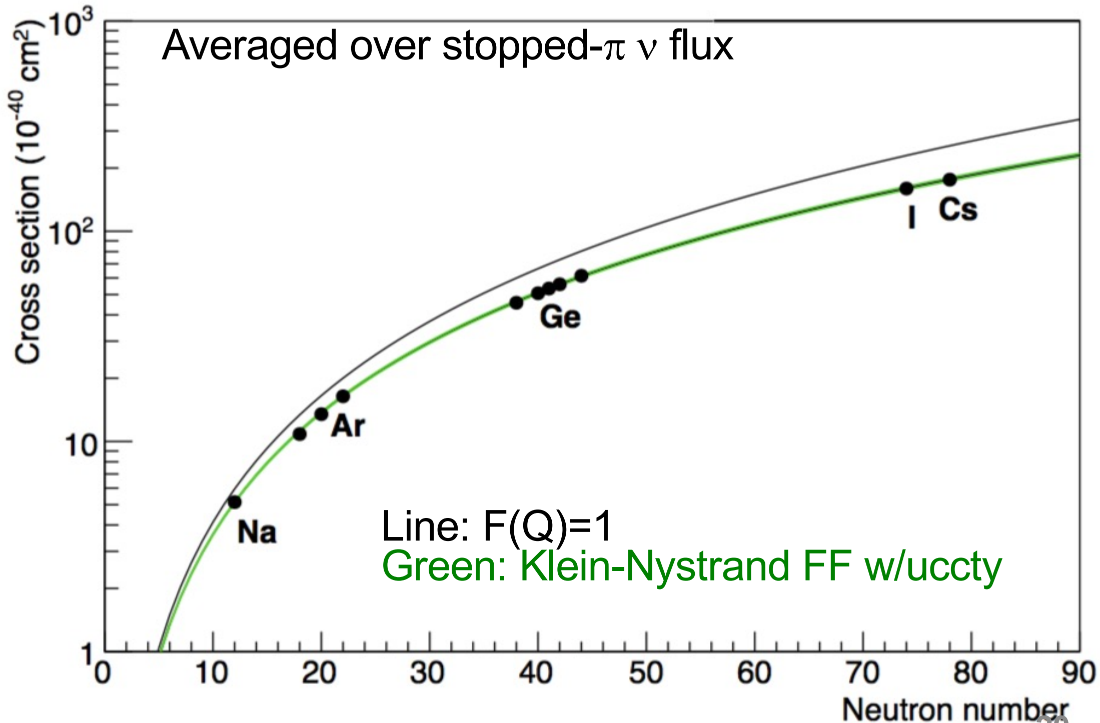
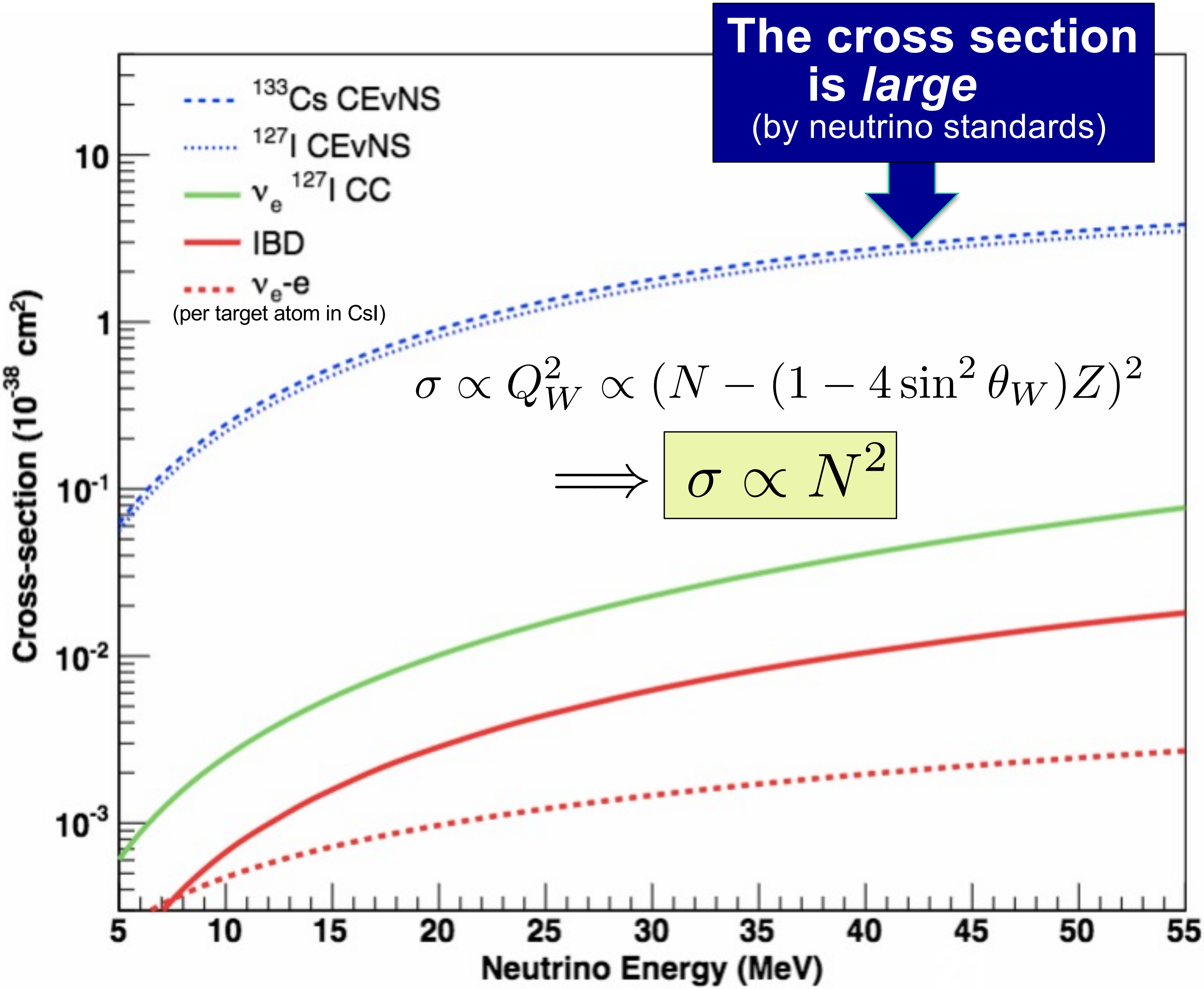
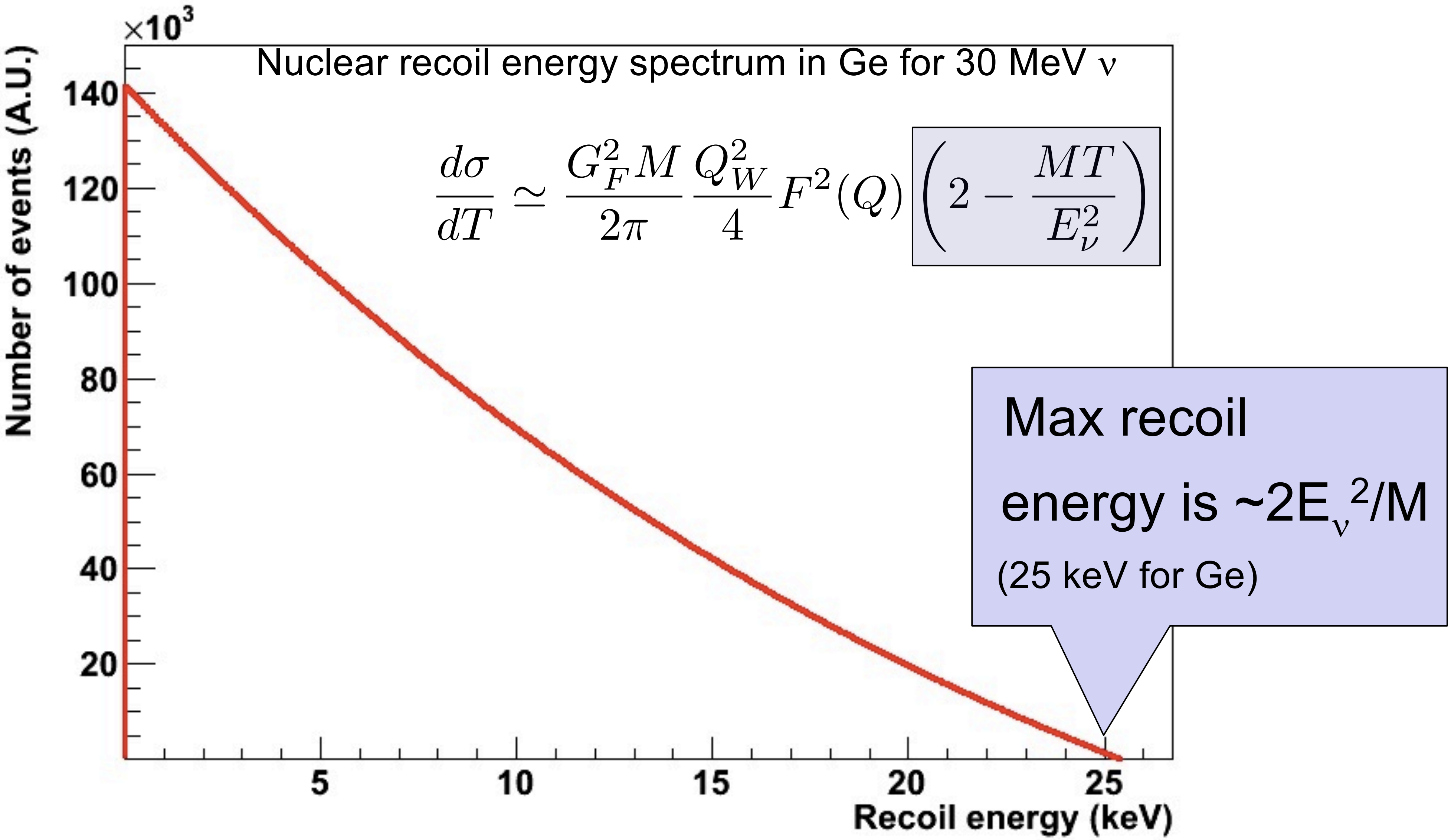
$$\Rightarrow \boxed{\frac{d\sigma}{dT} \propto N^2}$$





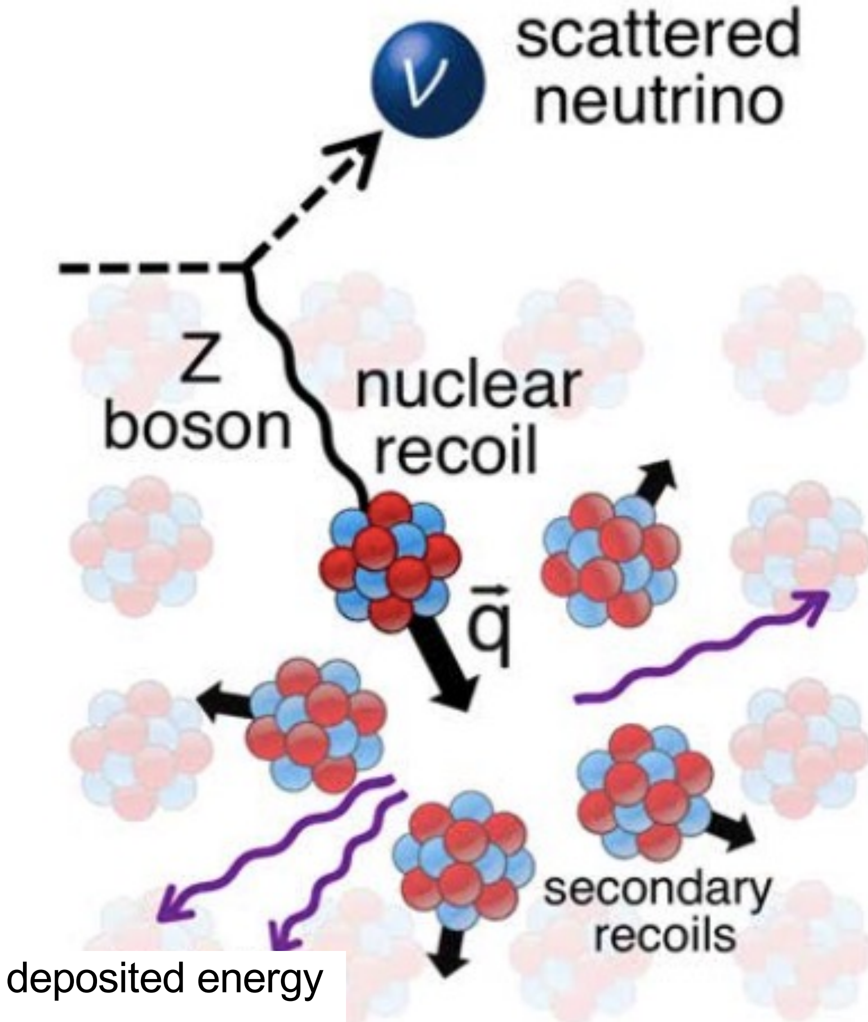
The cross section is *large*
(by neutrino standards)
^{133}Cs CEvNS
^{127}I CEvNS
ν_e ^{127}I CC
IBD
ν_e-e
(per target atom in CsI)
$\sigma \propto Q_W^2 \propto (N - (1 - 4\sin^2\theta_W)Z)^2$
$\Longrightarrow \boxed{\sigma \propto N^2}$
Cross-section (10^{-38} cm^2)
10
1
10^{-1}
10^{-2}
10^{-3}
5
10
15
20
25
30
35
40
45
50
55
Neutrino Energy (MeV)

Large cross section (by neutrino standards) but hard to observe due to **tiny nuclear recoil energies:**



The only experimental signature:

tiny energy deposited by nuclear recoils in the target material



➔ **WIMP dark matter detectors** developed over the last ~decade are sensitive to ~ keV to 10's of keV recoils

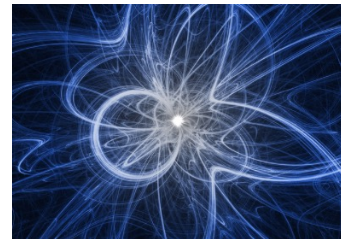
CEvNS: what's it good for?

1. So
2. Many
3. Things

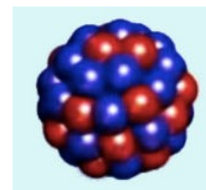
!

(not a complete list!)

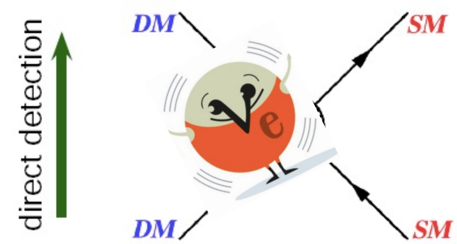
CEvNS as a **signal**
for signatures of *new physics*



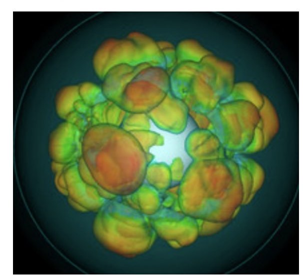
CEvNS as a **signal**
for understanding of "old" physics



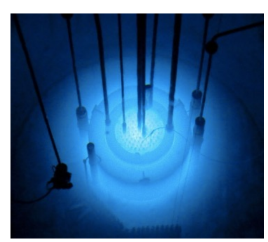
CEvNS as a **background**
for signatures of new physics



CEvNS as a **signal** for *astrophysics*



CEvNS as a **practical tool**



CEvNS: what's it good for?

! (not a complete list!)

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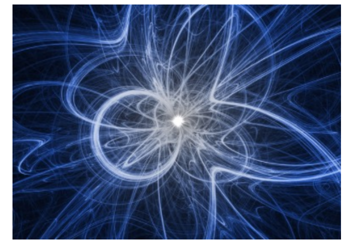

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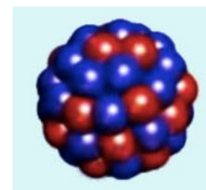

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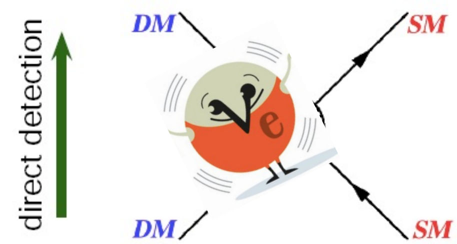

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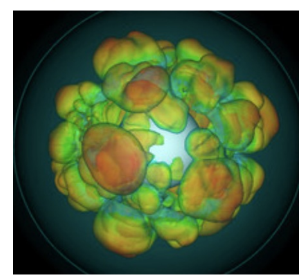

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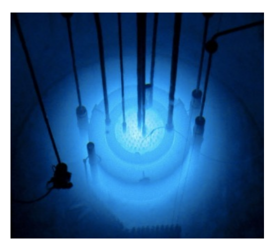

The cross section is cleanly predicted in the Standard Model

$$\frac{d\sigma}{dT} = \frac{G_F^2 M}{\pi} F^2(Q) \left[(G_V + G_A)^2 + (G_V - G_A)^2 \left(1 - \frac{T}{E_\nu}\right)^2 - (G_V^2 - G_A^2) \frac{MT}{E_\nu^2} \right]$$

E_ν: neutrino energy
T: nuclear recoil energy
M: nuclear mass
$Q = \sqrt{(2MT)}$: momentum transfer

G_V, G_A: SM weak parameters

vector $G_V = g_V^p Z + g_V^n N,$ ← dominates

axial $G_A = g_A^p (Z_+ - Z_-) + g_A^n (N_+ - N_-)$ ← small for most nuclei, zero for spin-zero

$$\boxed{\begin{aligned} g_V^p &= 0.0298 \\ g_V^n &= -0.5117 \\ g_A^p &= 0.4955 \\ g_A^n &= -0.5121. \end{aligned}}$$

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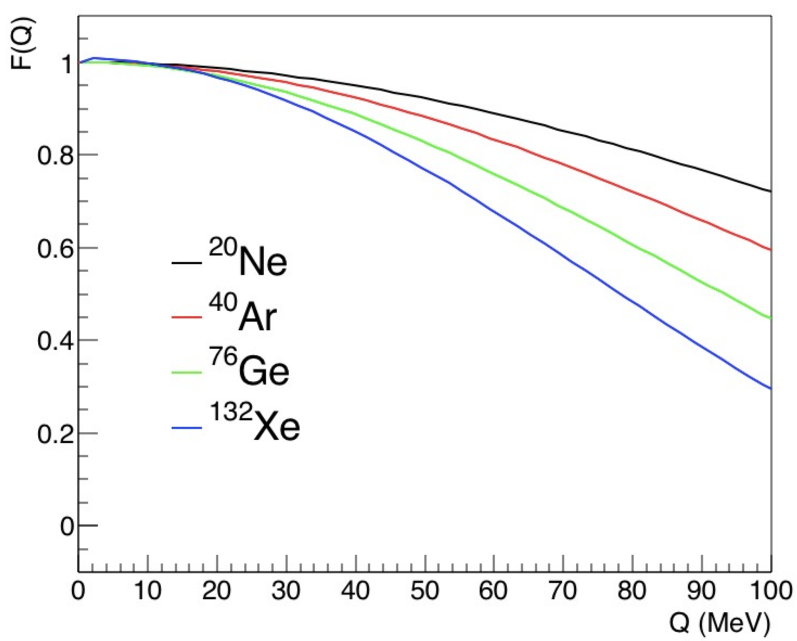
F(Q): nuclear **form factor**, <~5% uncertainty on event rate

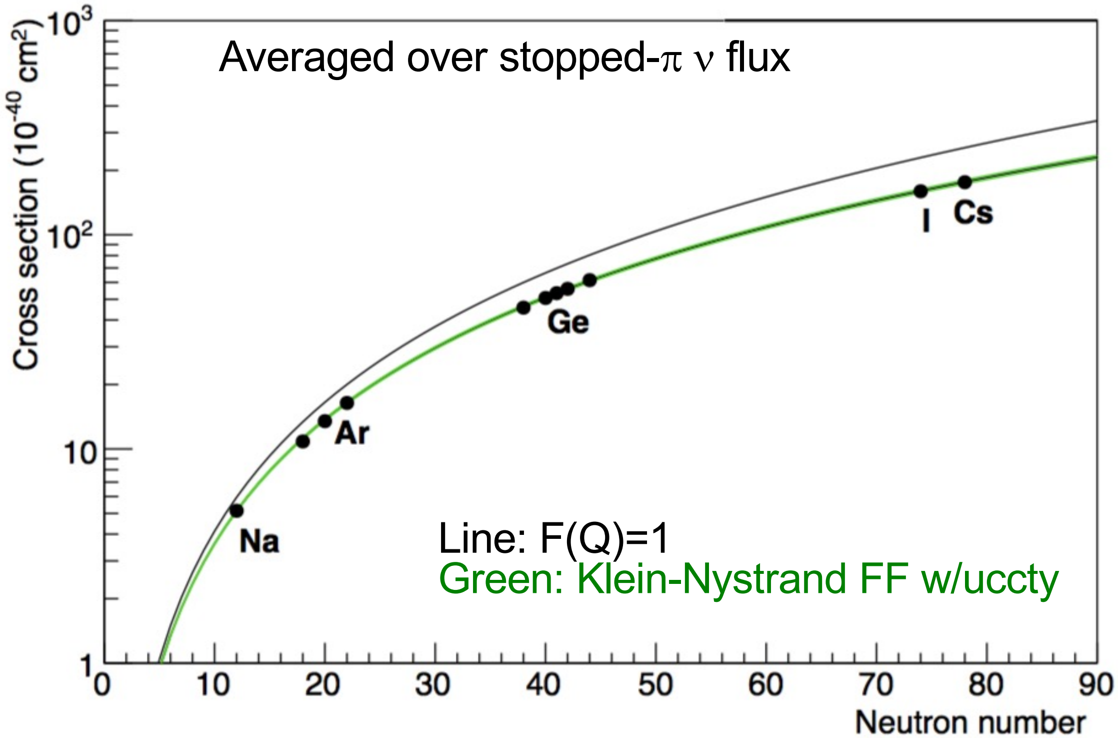


form factor suppresses cross section at large Q

The CEvNS rate is a clean SM prediction

$$\frac{d\sigma}{dT} = \frac{G_F^2 M}{2\pi} \frac{Q_W^2}{4} F^2(Q) \left(2 - \frac{MT}{E_\nu^2}\right)$$

small nuclear uncertainties



A deviation from α N^2 prediction can be a signature of beyond-the-SM physics

Searching for BSM Physics with CEvNS

A first example: simple counting to constrain **non-standard interactions (NSI)** of neutrinos with quarks

Davidson et al., JHEP 0303:011 (2004)
Barranco et al., JHEP 0512:021 (2005)

"Model-independent" parameterization

$$\mathcal{L}_{\nu H}^{NSI} = -\frac{G_F}{\sqrt{2}} \sum_{\substack{q=u,d \\ \alpha,\beta=e,\mu,\tau}} [\bar{\nu}_\alpha \gamma^\mu (1-\gamma^5) \nu_\beta] \times (\varepsilon_{\alpha\beta}^{qL} [\bar{q}\gamma_\mu (1-\gamma^5) q] + \varepsilon_{\alpha\beta}^{qR} [\bar{q}\gamma_\mu (1+\gamma^5) q])$$

ε's parameterize new interactions

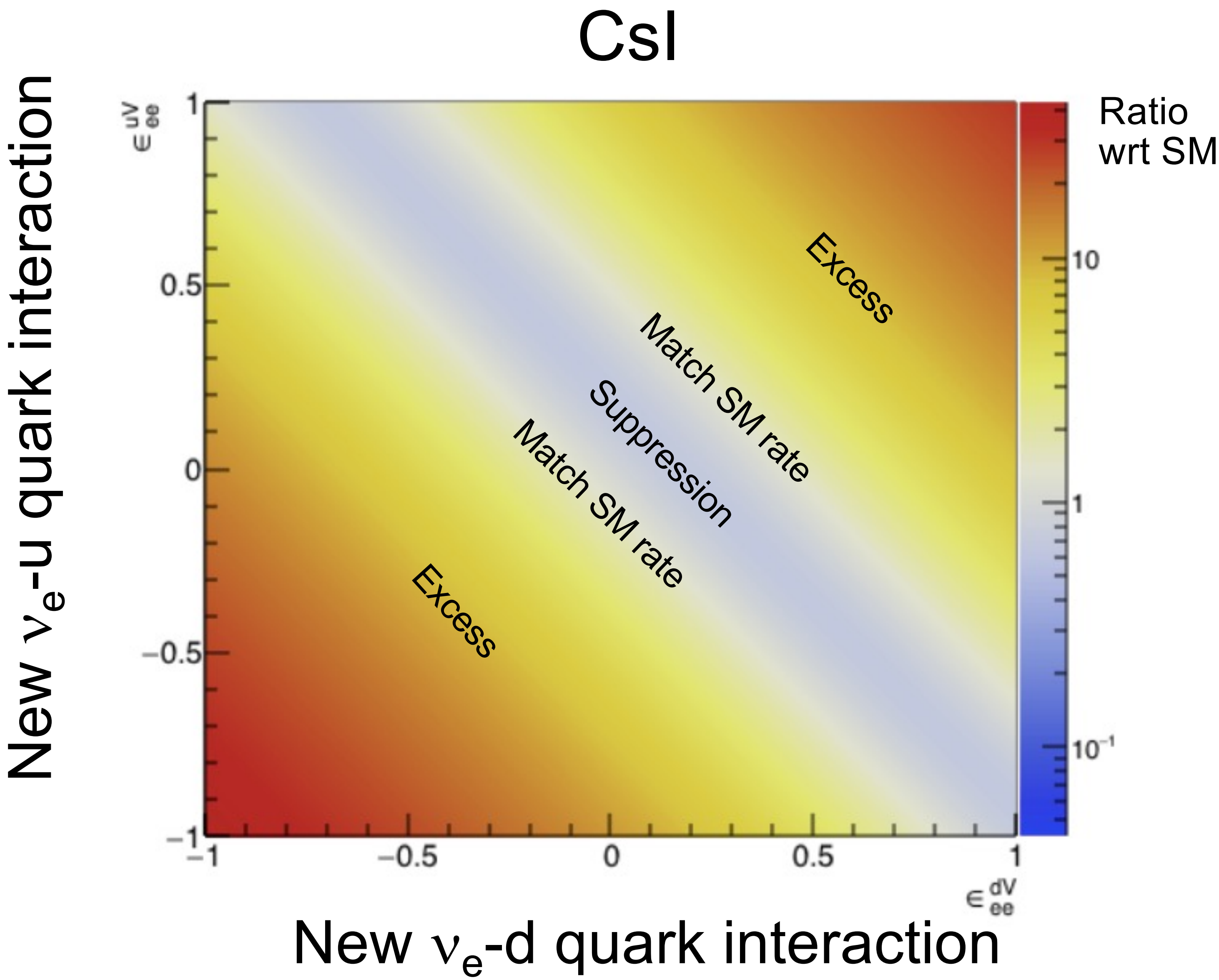
"Non-Universal": ε_{ee}, $\varepsilon_{\mu\mu}$, $\varepsilon_{\tau\tau}$

Flavor-changing: $\varepsilon_{\alpha\beta}$, where $\alpha \neq \beta$

$\Rightarrow$ some are quite poorly constrained (~unity allowed)

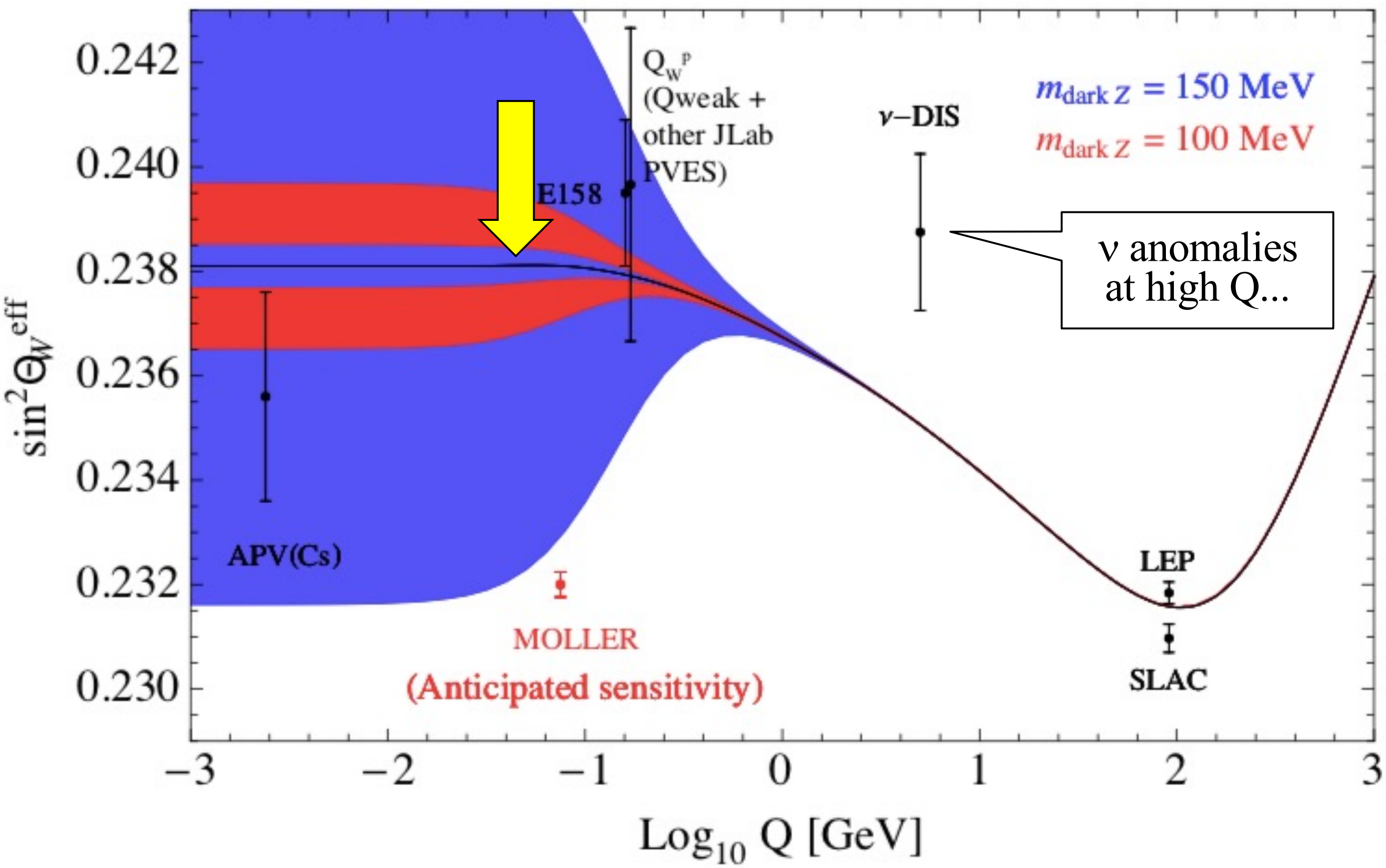
Signatures of **Beyond-the-Standard-Model Physics**

Look for a CEvNS **excess** or **deficit** wrt SM expectation



A new method for measuring the **weak mixing angle** at Q~0.04 GeV

$$\sigma \propto Q_W^2 \propto (N - (1 - 4\sin^2\theta_W)Z)^2$$



MOLLER collaboration, arXiv:1411.4088

Other new physics results in a *distortion of the recoil spectrum* (Q dependence)

BSM Light Mediators

SM weak charge

$$Q^2_{\alpha,\mathrm{SM}} = (Z g_p^V + N g_n^V)^2$$

Effective weak charge in presence of light vector mediator Z'

$$Q^2_{\alpha,\mathrm{NSI}} = \left[Z\left(g_p^V + \frac{3g^2}{2\sqrt{2}G_F(Q^2 + M^2_{Z'})} \right) + N\left(g_n^V + \frac{3g^2}{2\sqrt{2}G_F(Q^2 + M^2_{Z'})} \right) \right]^2$$

specific to neutrinos and quarks

e.g. arXiv:1708.04255

Neutrino (Anomalous) Magnetic Moment

e.g. arXiv:1505.03202, 1711.09773

$$\left(\frac{d\sigma}{dT}\right)_m = \frac{\pi \alpha^2 \mu_\nu^2 Z^2}{m_e^2} \left(\frac{1 - T/E_\nu}{T} + \frac{T}{4E_\nu^2} \right)$$

Specific ~1/T upturn at low recoil energy

Sterile Neutrino Oscillations

$$P^{\mathrm{SBL}}_{\nu_\alpha \to \nu_\alpha}(E_\nu) = 1 - \sin^2 2\theta_{\alpha\alpha} \sin^2\left(\frac{\Delta m^2_{41} L}{4E_\nu}\right)$$

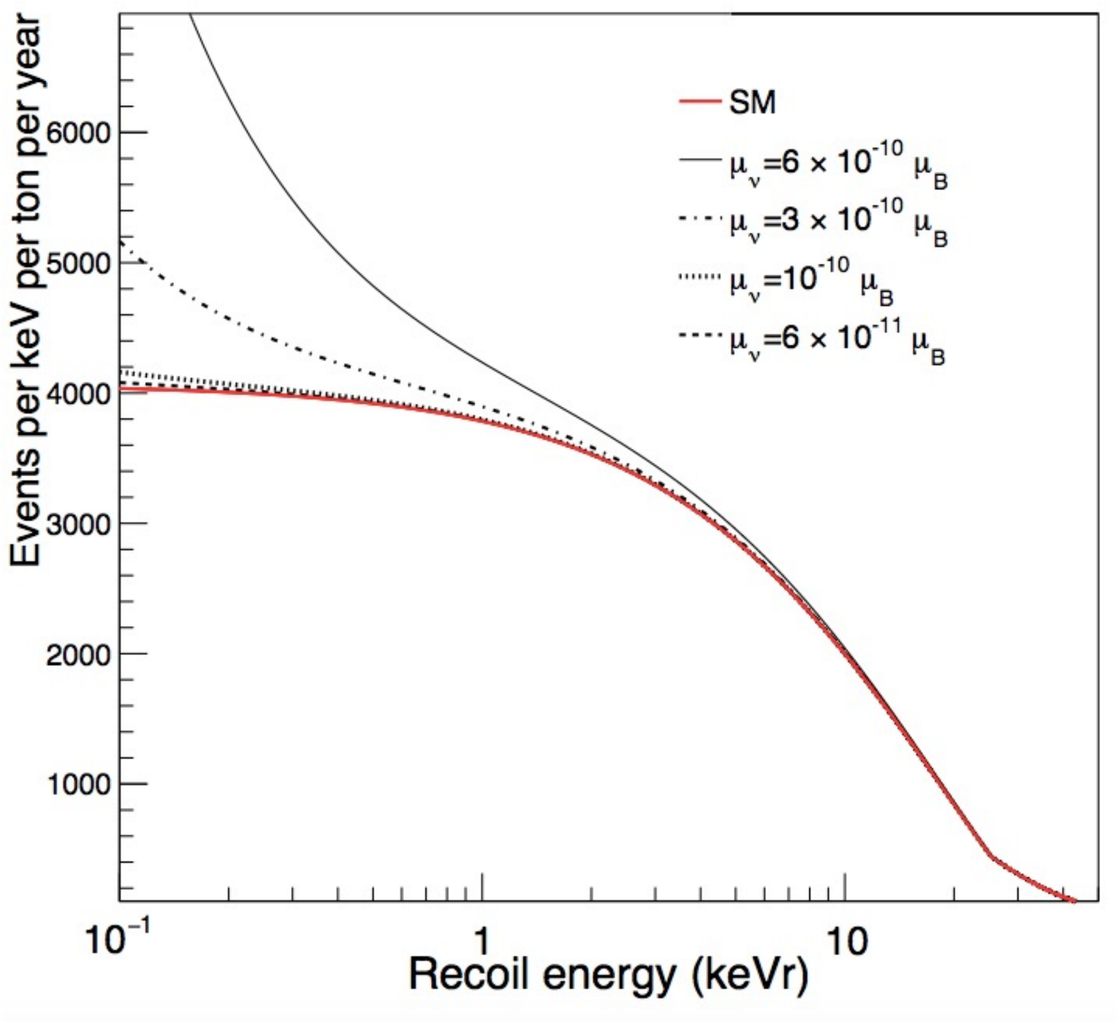
"True" disappearance with baseline-dependent Q distortion

e.g. arXiv: 1511.02834, 1711.09773, 1901.08094

Neutrino magnetic moment

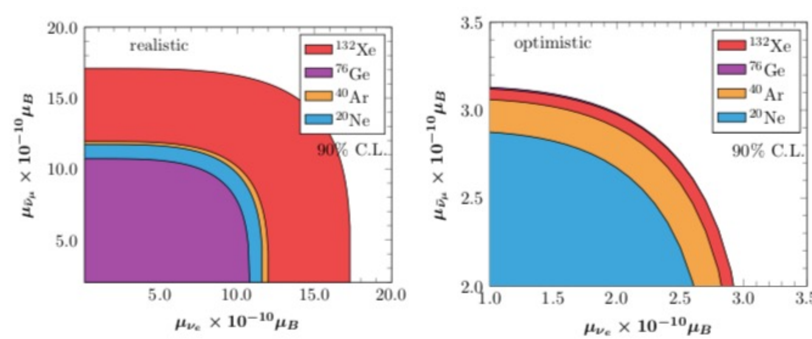
Signature is **distortion at low recoil energy T**

$$\left(\frac{d\sigma}{dT}\right)_m = \frac{\pi\alpha^2\mu_\nu^2 Z^2}{m_e^2}\left(\frac{1-T/E_\nu}{T}+\frac{T}{4E_\nu^2}\right)$$



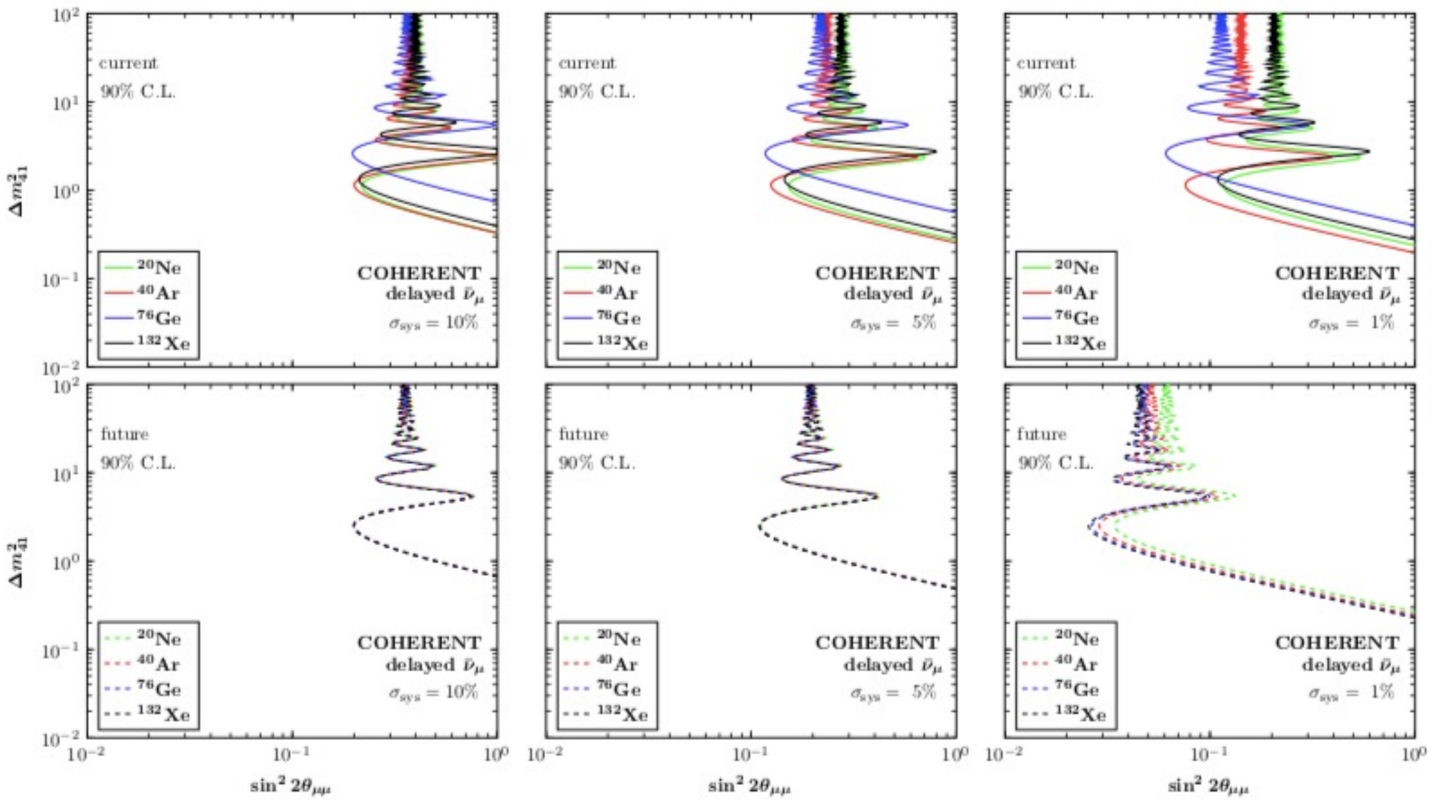
➔requires very low energy threshold (i.e., Ge)

See also Kosmas et al., arXiv:1505.03202, 1711.09773



Another example: **sterile neutrino oscillations**

- CEvNS is NC and doesn't care about the flavor;
  disappearance is "true" disappearance
- Some neutrino spectral info in the recoil spectrum
- Can cancel some systematics with multiple identical, or movable detectors



In this work, the calculation is performed for two cases corresponding to (i) the "current" configuration: a (^{20}Ne, ^{40}Ar, ^{76}Ge, ^{132}Xe) target with mass (391, 456, 100, 100) kg located at (46, 46, 20, 40) m from the source with energy threshold of (30, 20, 10, 8) keV$_{nr}$ and a running time of 2.4×10^7s, and (ii) the "future" configuration: 1 ton of detector mass located at 20 m from the source with energy threshold 1 keV$_{nr}$ and 1 year of data taking time (see e.g Ref. [7]).

Even 100 kg of Ge is expensive/challenging, but multitons of noble liquid is entirely thinkable

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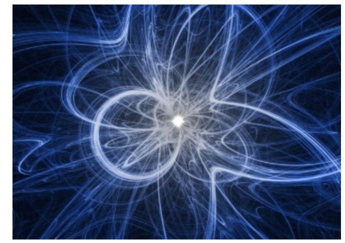

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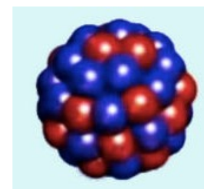
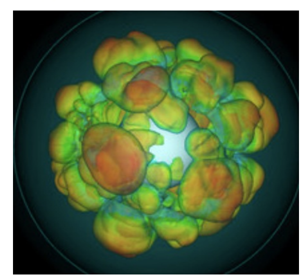
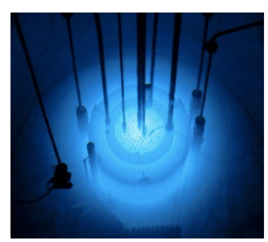


CEvNS as a **background**
for signatures of new physics

CEvNS as a **signal** for *astrophysics*



CEvNS as a **practical tool**



What can we learn about **nuclear physics** with CEvNS?

Nuclear neutron form factor from neutrino–nucleus coherent elastic scattering

P S Amanik and G C McLaughlin

Department of Physics, North Carolina State University, Raleigh, NC 27695-8202, USA

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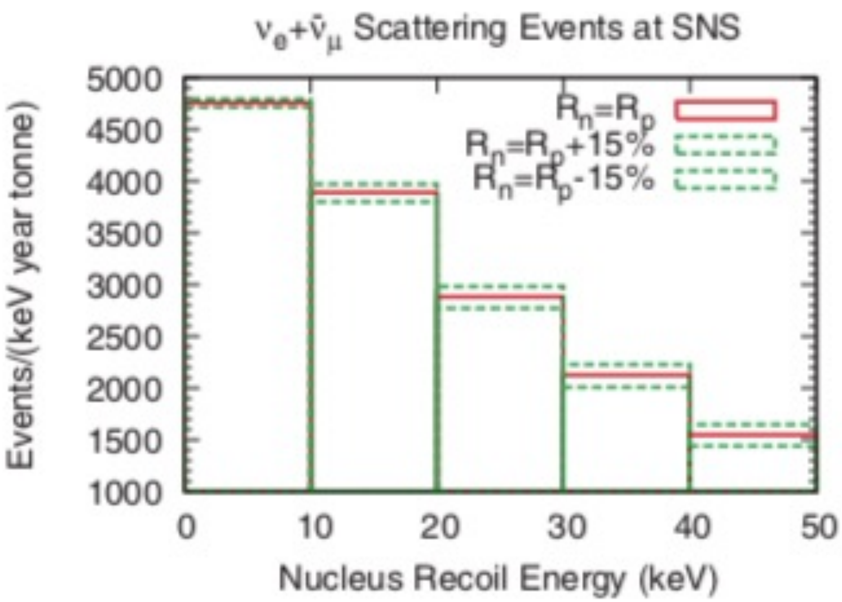
Abstract
We point out that there is potential to study the nuclear neutron form factor through neutrino nucleus coherent elastic scattering. We determine numbers of events for various scenarios in a liquid noble nuclear recoil detector at a stopped pion neutrino source.

Neutrino-nucleus coherent scattering as a probe of neutron density distributions

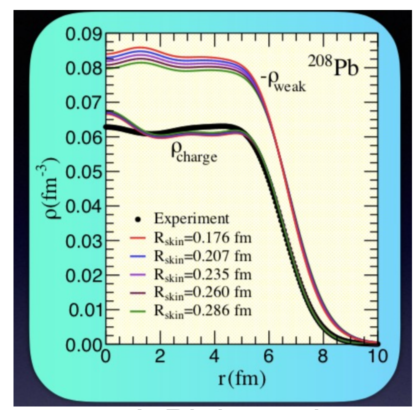
Kelly Patton[1,*] Jonathan Engel[2,†] Gail C. McLaughlin[1,‡] and Nicolas Schunck[3,§]
[1] *Physics Department, North Carolina State University, Raleigh, North Carolina 27695, USA*
[2] *Department of Physics and Astronomy, University of North Carolina, Chapel Hill, North Carolina 27599, USA*
[3] *Physics Division, Lawrence Livermore Laboratory, Livermore, California 94551 USA*
(Dated: July 4, 2012)

Neutrino-nucleus coherent elastic scattering provides a theoretically appealing way to measure the neutron part of nuclear form factors. Using an expansion of form factors into moments, we show that neutrinos from stopped pions can probe not only the second moment of the form factor (the neutron radius) but also the fourth moment. Using simple Monte Carlo techniques for argon, germanium, and xenon detectors of 3.5 tonnes, 1.5 tonnes, and 300 kg, respectively, we show that the neutron radii can be found with an uncertainty of a few percent when near a neutrino flux of 3×10^7 neutrinos/cm^2/s. If the normalization of the neutrino flux is known independently, one can determine the moments accurately enough to discriminate among the predictions of various nuclear energy functionals.



Observable is **recoil spectrum shape**

Neutron radius and "skin" (R_n-R_p) relevant for understanding of neutron stars

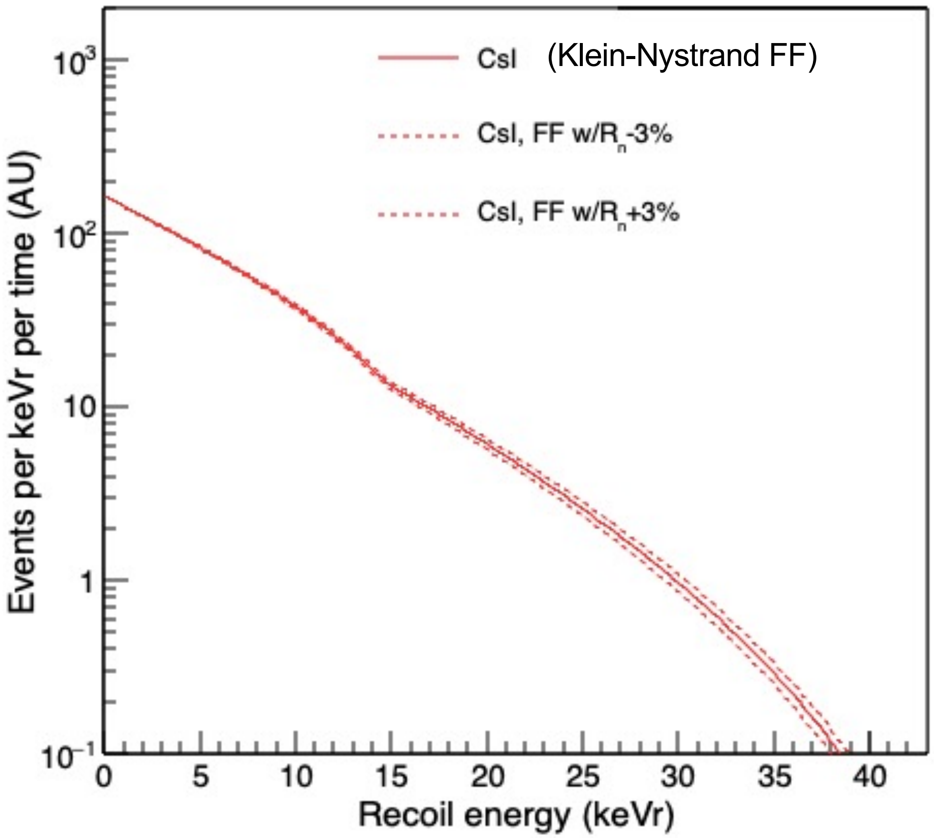


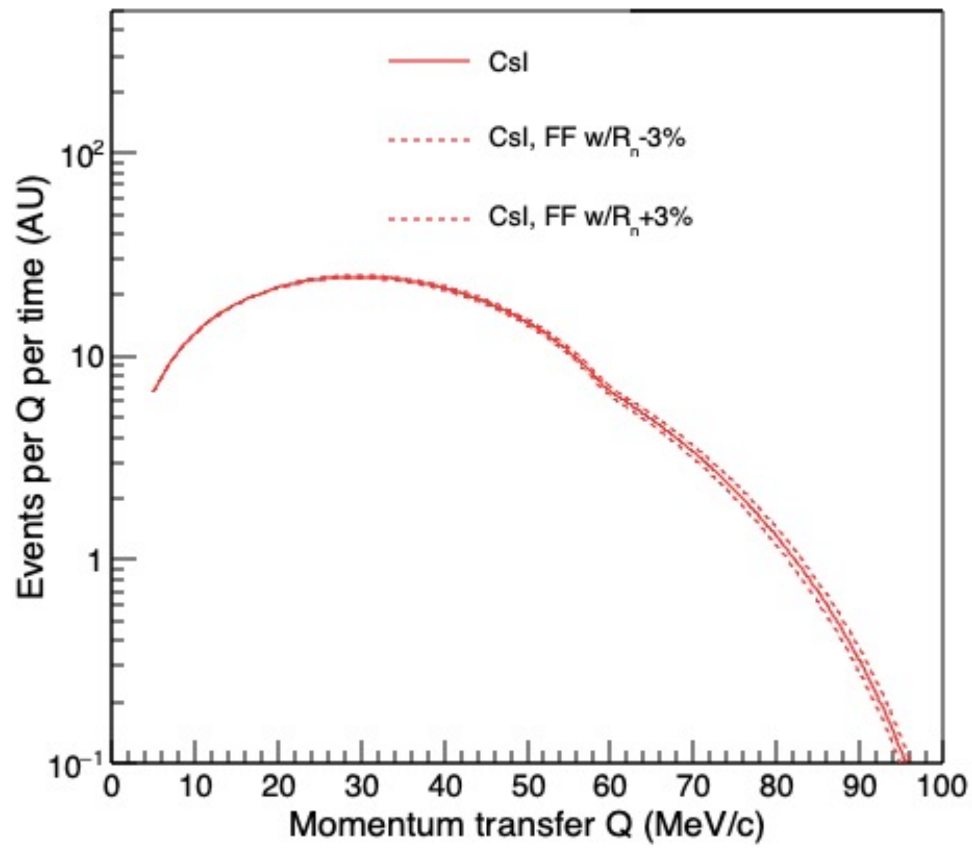
J. Piekarewicz

Effect of form-factor *uncertainty* on the recoil spectrum: estimate as R_n +/- 3%

$$\frac{d\sigma}{dT} = \frac{G_F^2 M}{\pi} F^2(Q) \left[(G_V + G_A)^2 + (G_V - G_A)^2 \left(1 - \frac{T}{E_\nu}\right)^2 - (G_V^2 - G_A^2)\frac{MT}{E_\nu^2} \right]$$

Stopped-π spectrum





At current level of experimental precision, form factor uncertainty is small effect

So: if you are hunting for BSM physics as a distortion of the recoil spectrum

. . . uncertainties in the form factor are a nuisance!

There are degeneracies in the observables between "old" (but still mysterious) physics

and "new" physics

We will need to think carefully about how to disentangle these effects and understand uncertainties, for the longer term

[See also: D. Aristizabal Sierra et al. arXiv:1902.07398, recent INT workshop "Weak Elastic Scattering with Nuclei"]

Summary of what we can get at experimentally

Observables:

Event rate
Recoil spectrum ($T=Q^2/2M$)
[In principle: scattering angle... hard]

Spectral shape systematics are hard!

Knowable/controllable parameters:

Neutrino flavor, via source, and timing

(reactor: ν_e-bar, stopped-π: ν_e, ν_μ-bar, ν_μ)

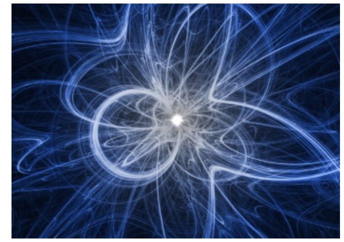
N, Z via nuclear target type
Baseline
Direction with respect to source

CEvNS: what's it good for?

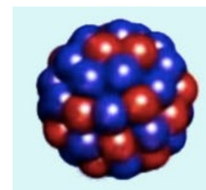
1 So
2 Many
3 Things
!

(not a complete list!)

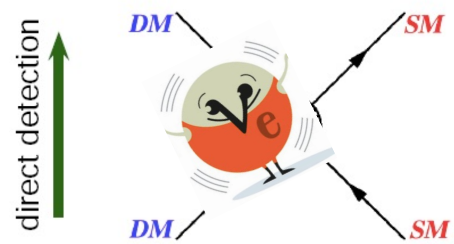
CEvNS as a **signal**
for signatures of *new physics*



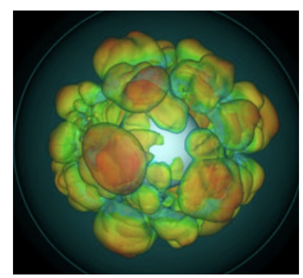
CEvNS as a **signal**
for understanding of "old" physics



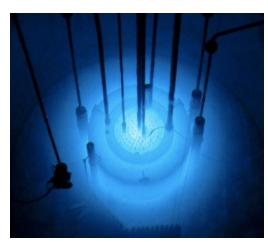
CEvNS as a **background**
for signatures of new physics (DM)



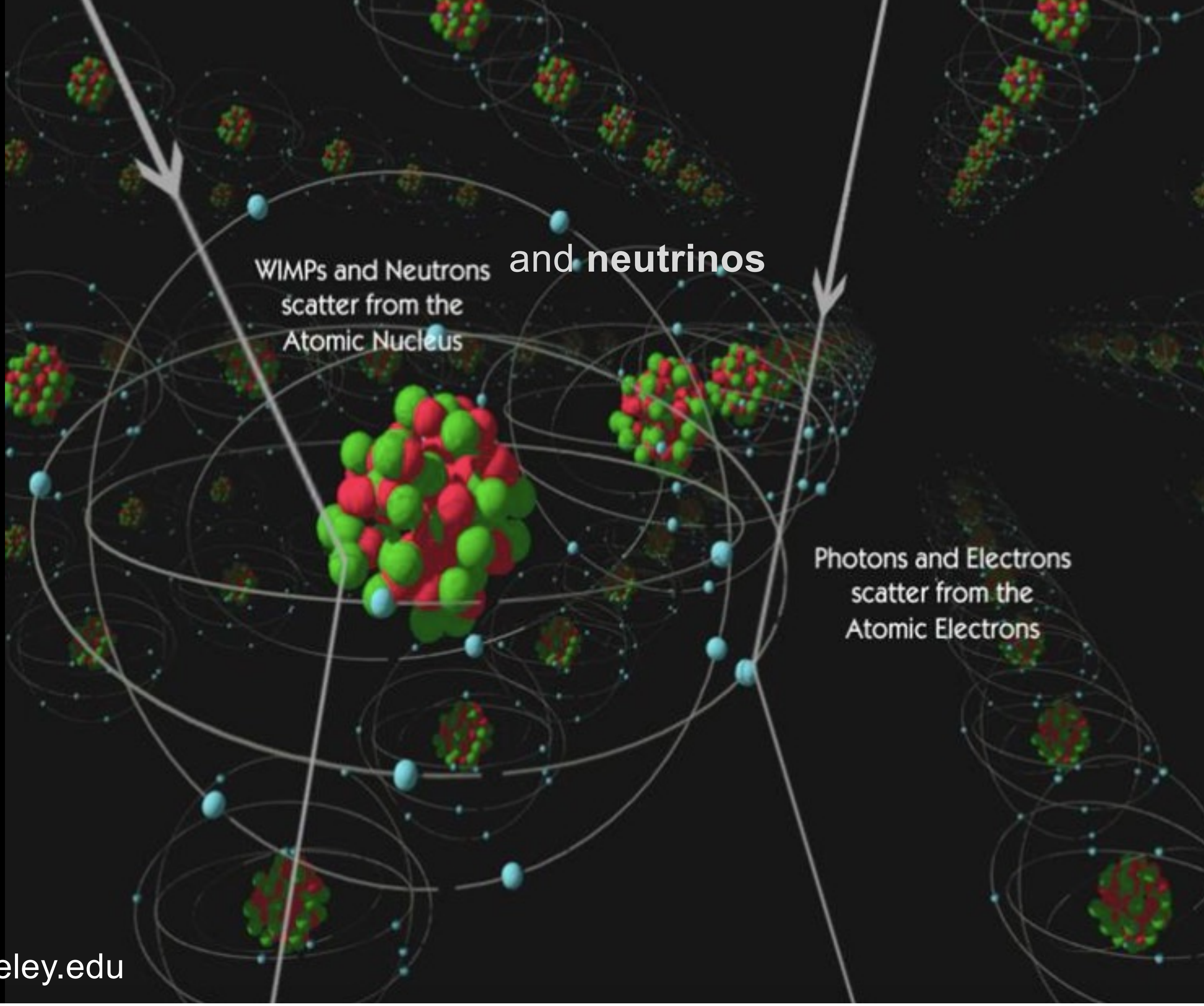
CEvNS as a **signal** for *astrophysics*



CEvNS as a **practical tool**

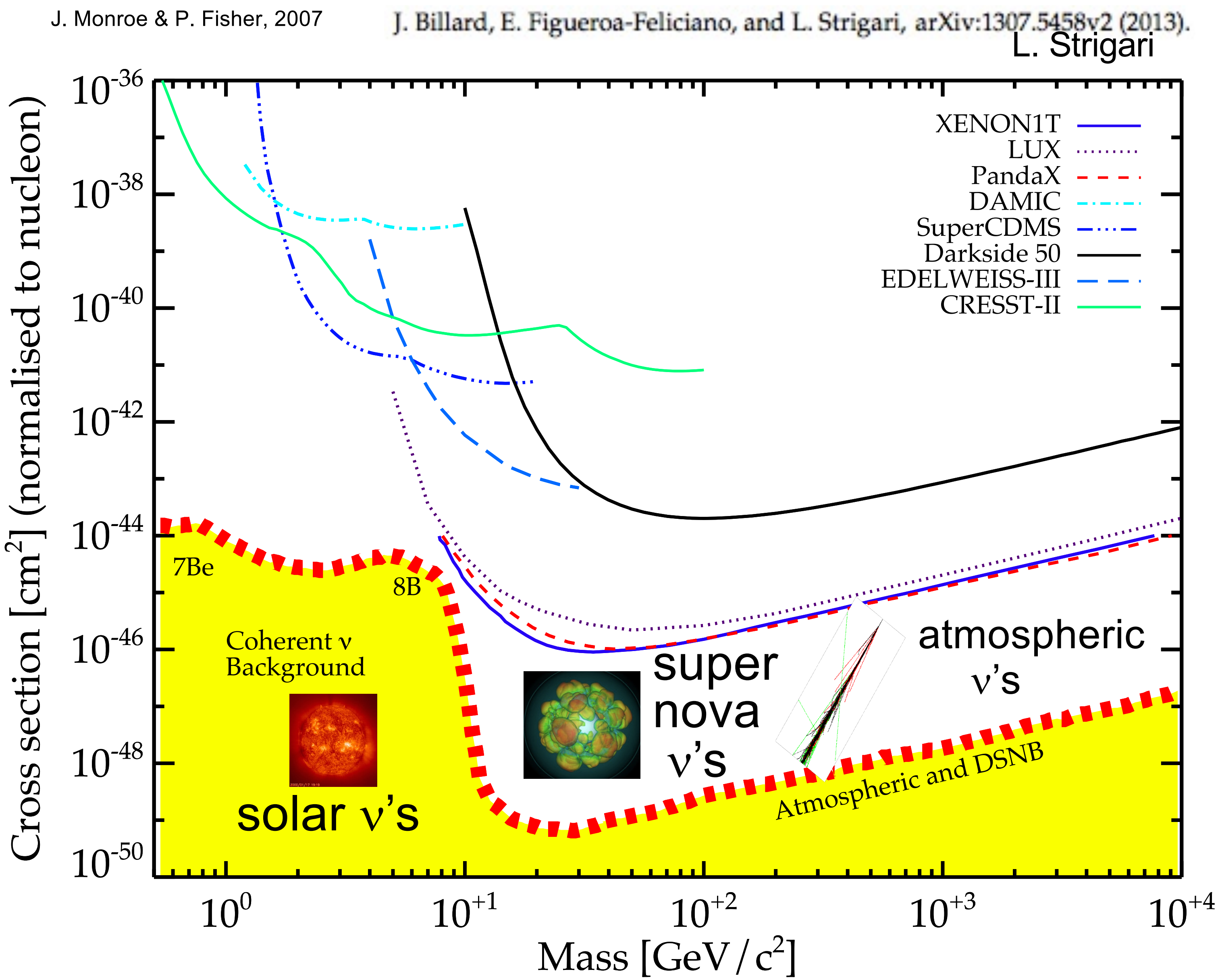


CEvNS from natural neutrinos creates ultimate background for direct DM search experiments



cdms.berkeley.edu

The so-called “neutrino floor” (**signal!**) for direct DM experiments



Light **accelerator-produced DM** direct detection possibilities

(CEvNS is *background*)

- "Vector portal": mixing of vector mediator with photons in π^0/η^0 decays
- "Leptophobic portal": new mediator coupling to baryons

$\Big\}$ decay product χ then makes nuclear recoil

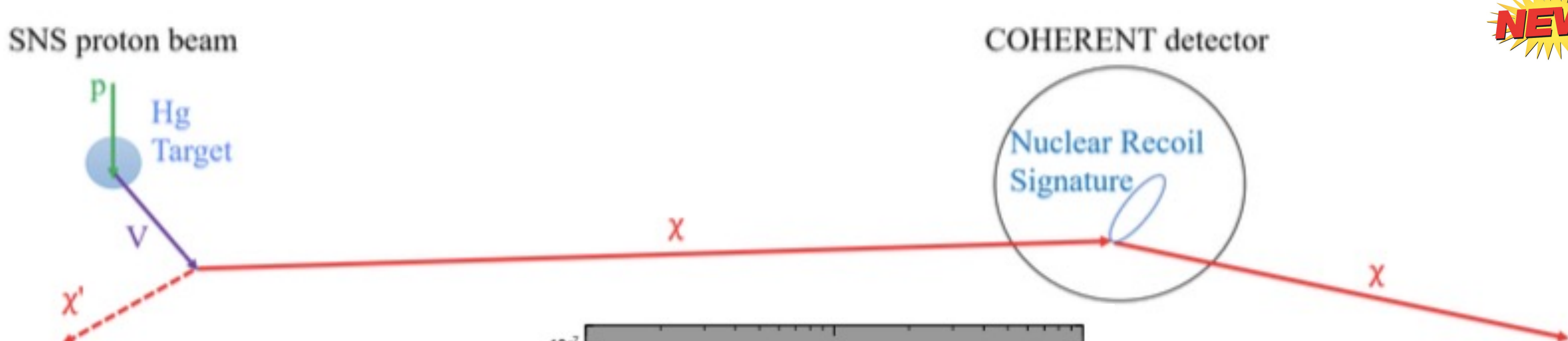
$$\pi^0 \longrightarrow \gamma + V^{(*)} \longrightarrow \gamma + \chi^\dagger + \chi$$

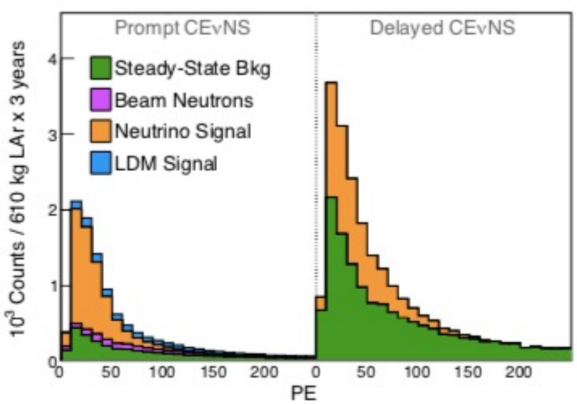
$$\pi^- + p \longrightarrow n + V^{(*)} \longrightarrow n + \chi^\dagger + \chi$$

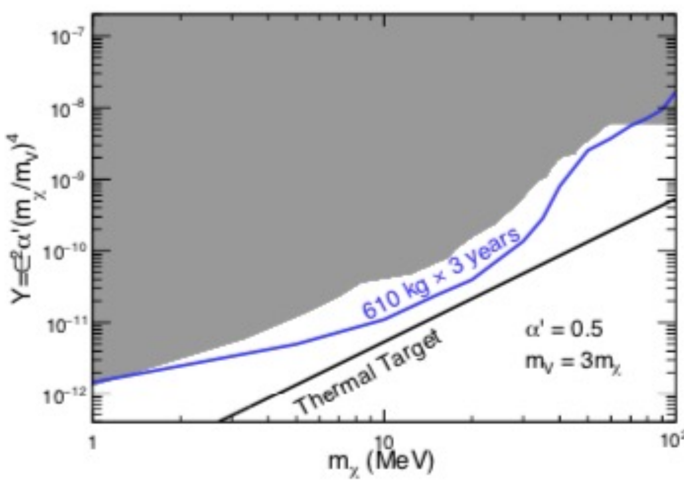
B. Batell et al., PRD 90 (2014)
P. de Niverville et al., PRD 95 (2017)

B. Dutta et al., arXiv:1906.10745
COHERENT, arXiv:1911.6422

NEW!





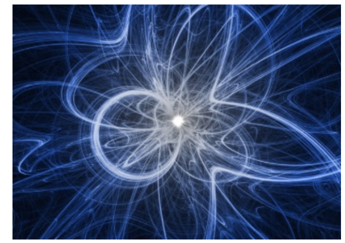


Expect *characteristic time, recoil energy, angle distribution* for DM vs CEvNS

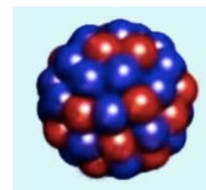
CEvNS: what's it good for?

(1) So (2) Many (3) Things **!** (not a complete list!)

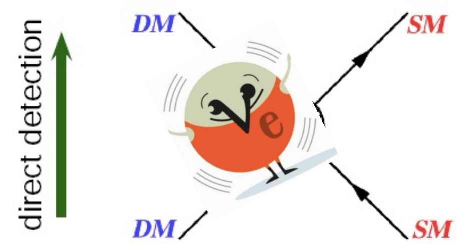
CEvNS as a **signal**
for signatures of *new physics*



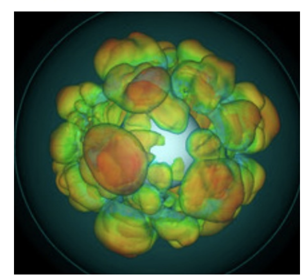
CEvNS as a **signal**
for understanding of "old" physics



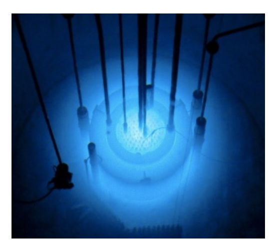
CEvNS as a **background**
for signatures of new physics (DM)



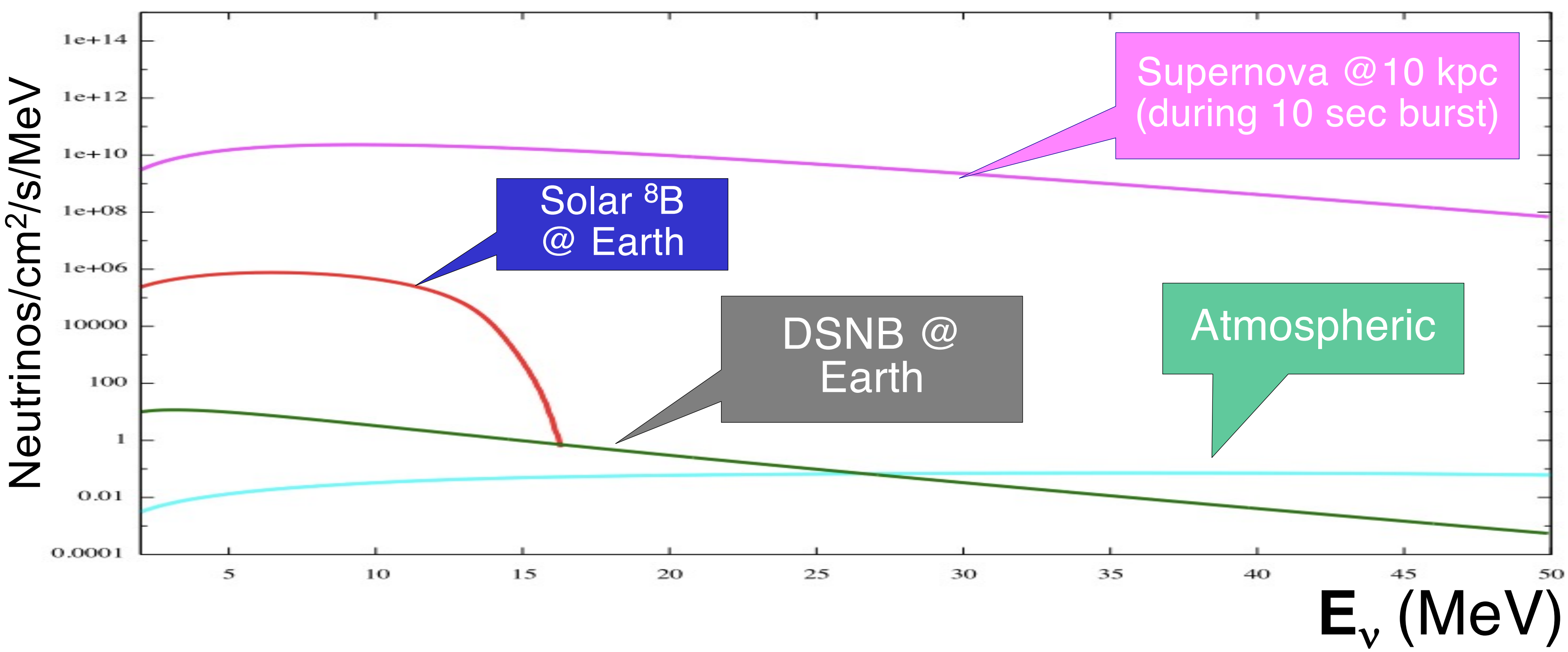
CEvNS as a **signal** for *astrophysics*



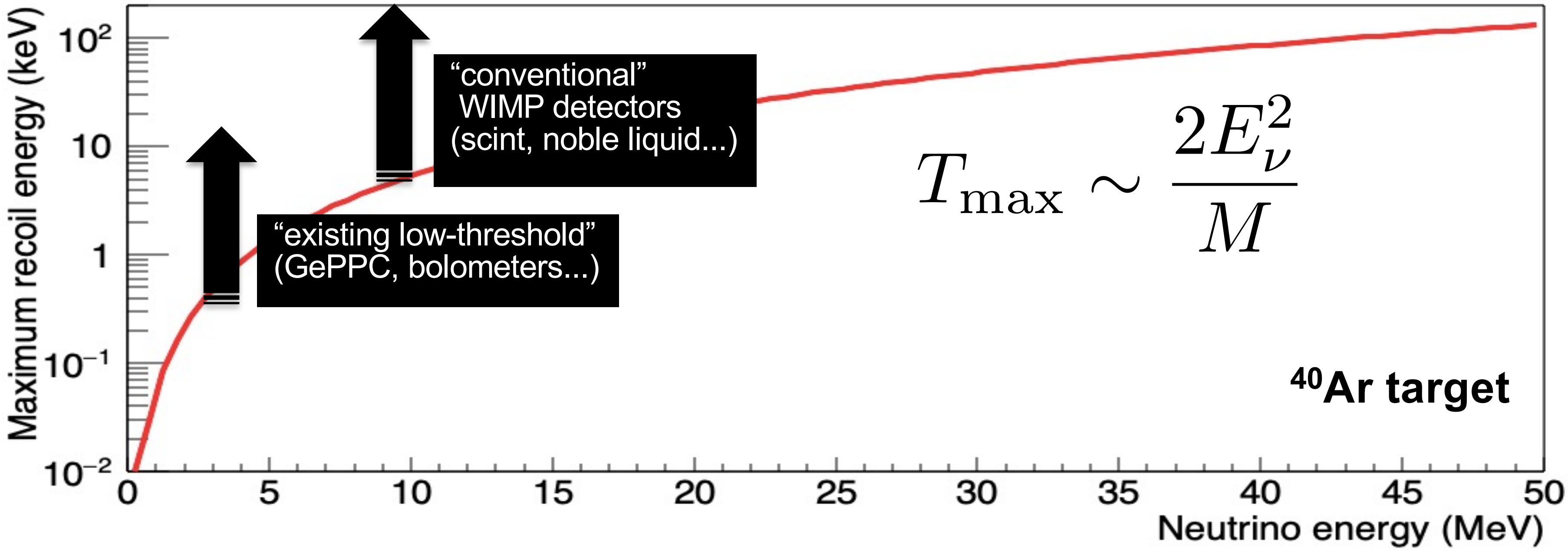
CEvNS as a **practical tool**

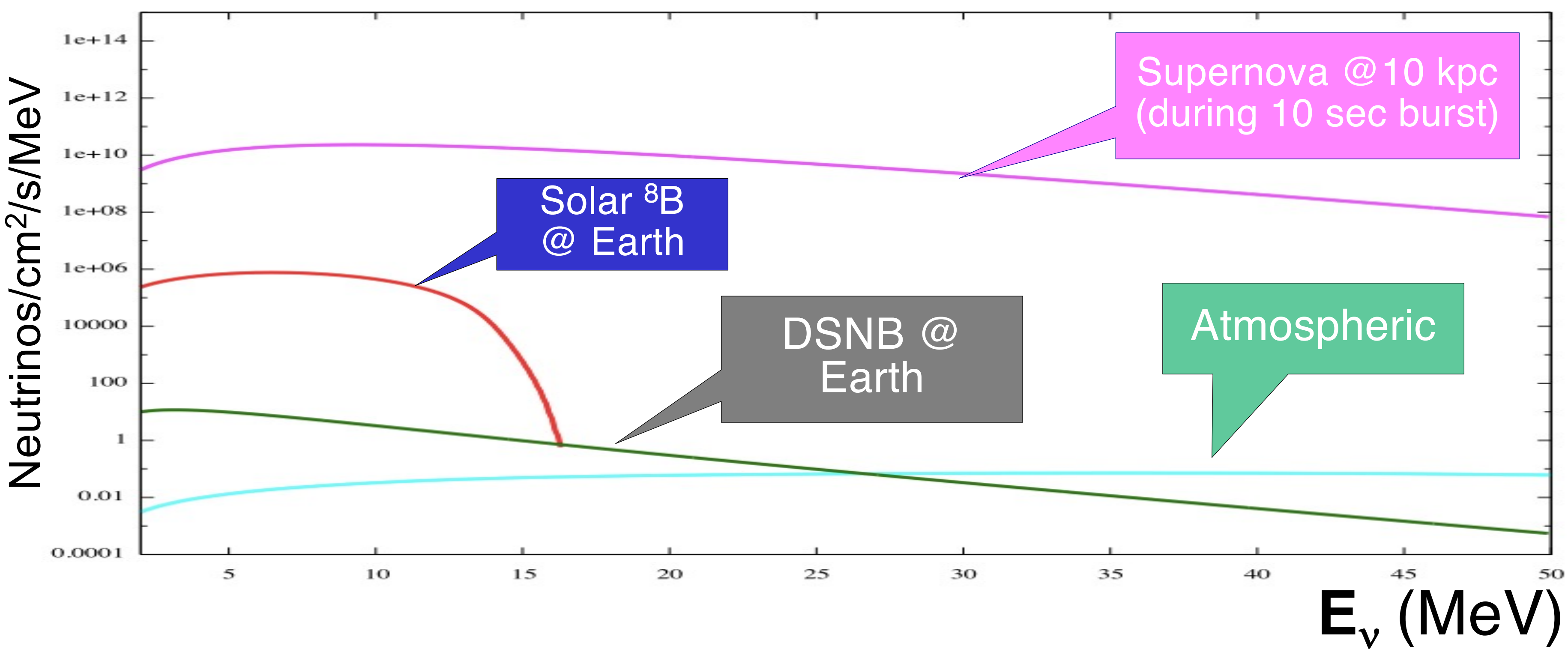


Natural neutrino fluxes



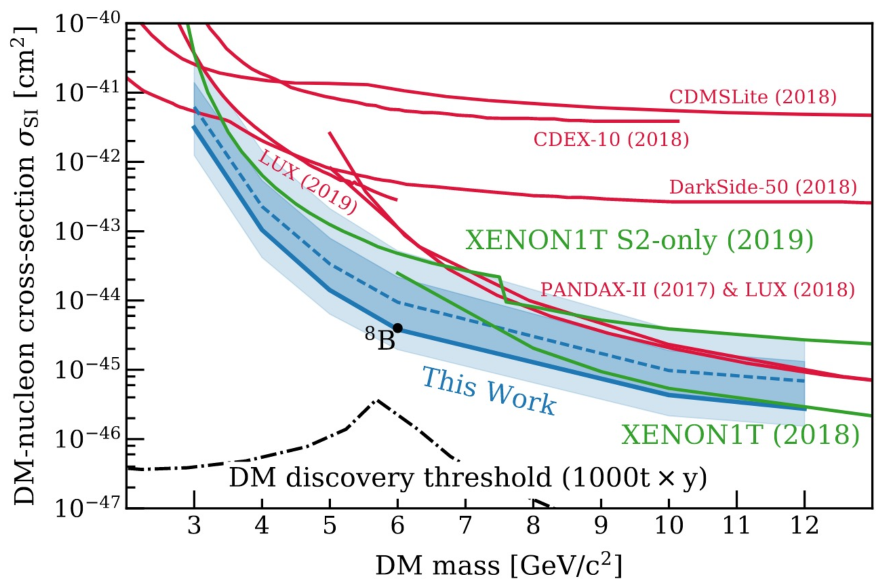
Natural neutrino fluxes



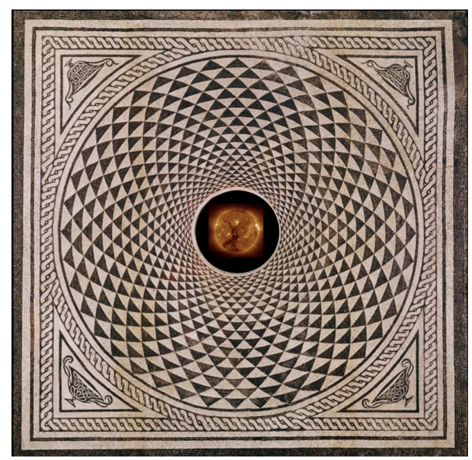


Search for CEvNS from solar neutrinos with the XENON-1T experiment

Phys.Rev.Lett. 126 (2021) 091301, arXiv: 2012.02846



Limit only so far
... but will eventually hit the floor...
sometimes there are interesting things to see if you look down...

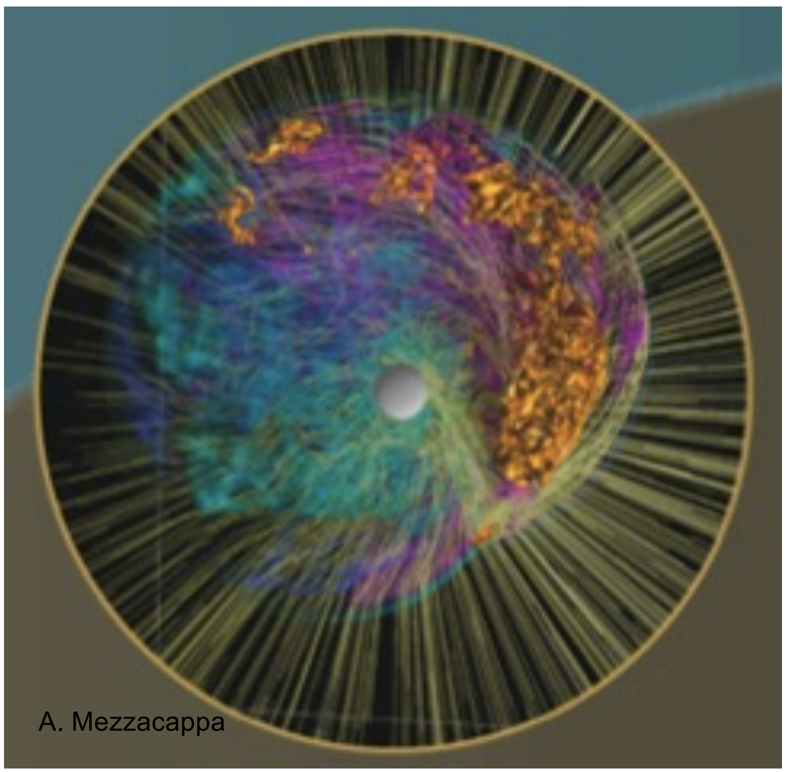


Neutrinos from core-collapse supernovae

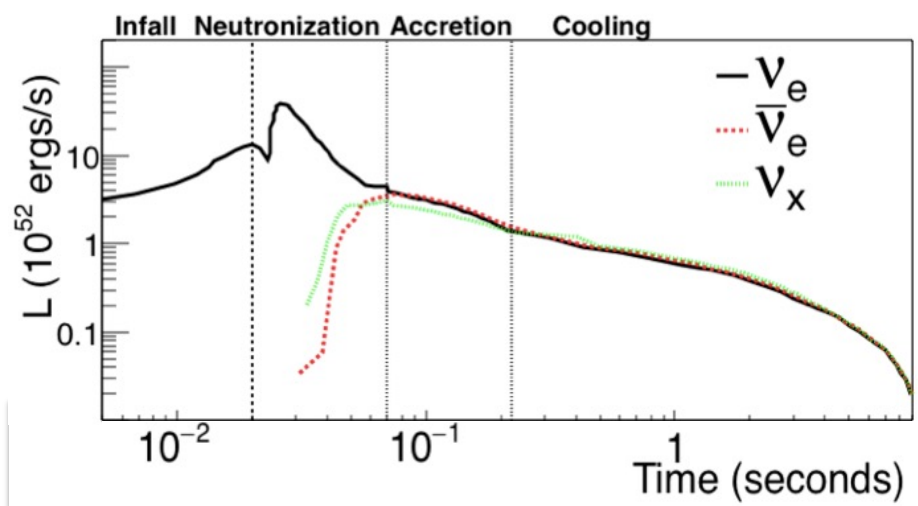
When a star's core collapses, ~99% of the gravitational binding energy of the proto-nstar goes into ν's of ***all flavors*** with ~tens-of-MeV energies

Energy *can* escape via ν's

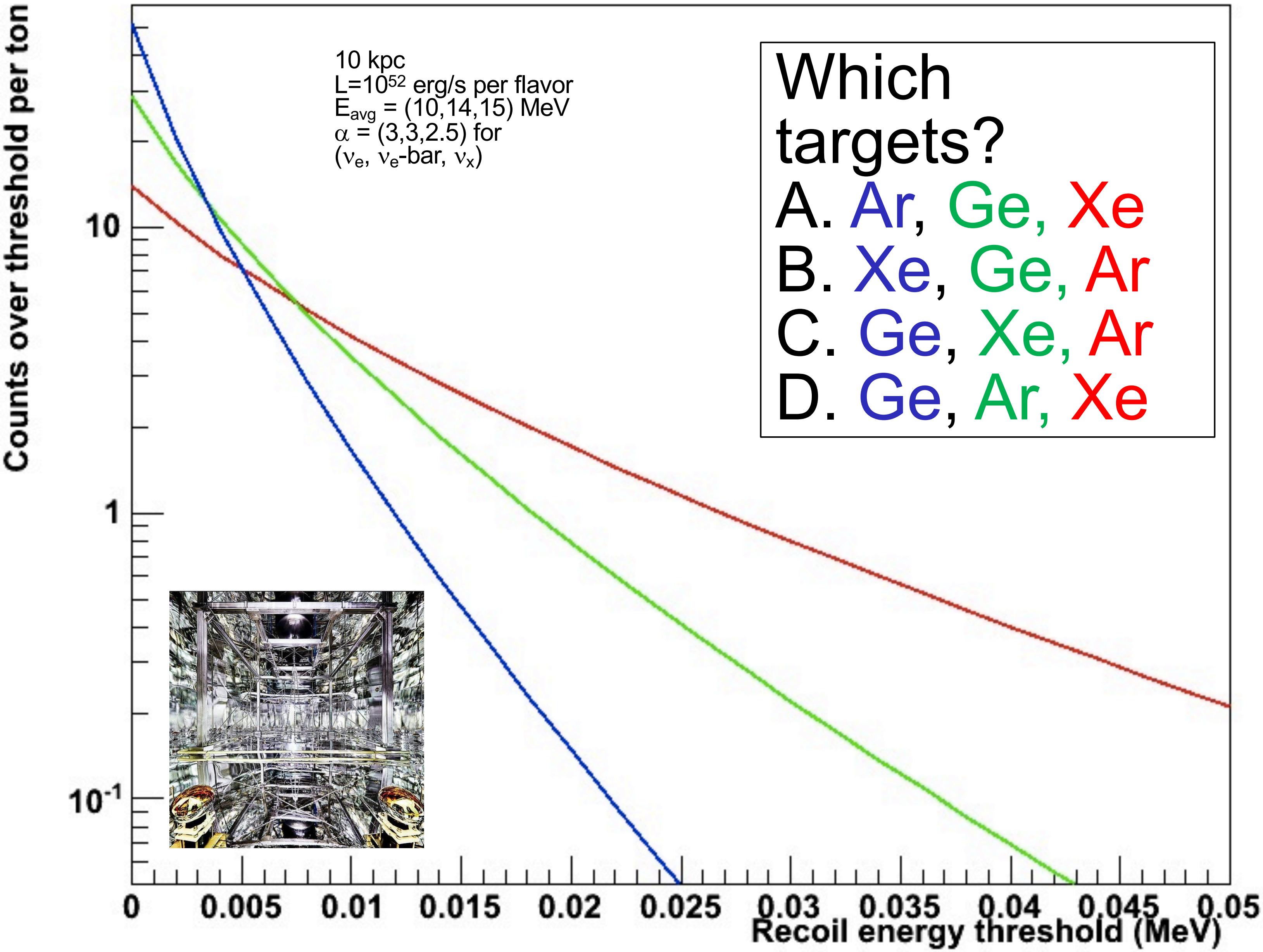
Mostly ν-$\bar{\nu}$ pairs from proto-nstar cooling



Timescale: *prompt* after core collapse, overall Δt~10's of seconds



Supernova neutrinos in tonne-scale DM detectors



Which targets?

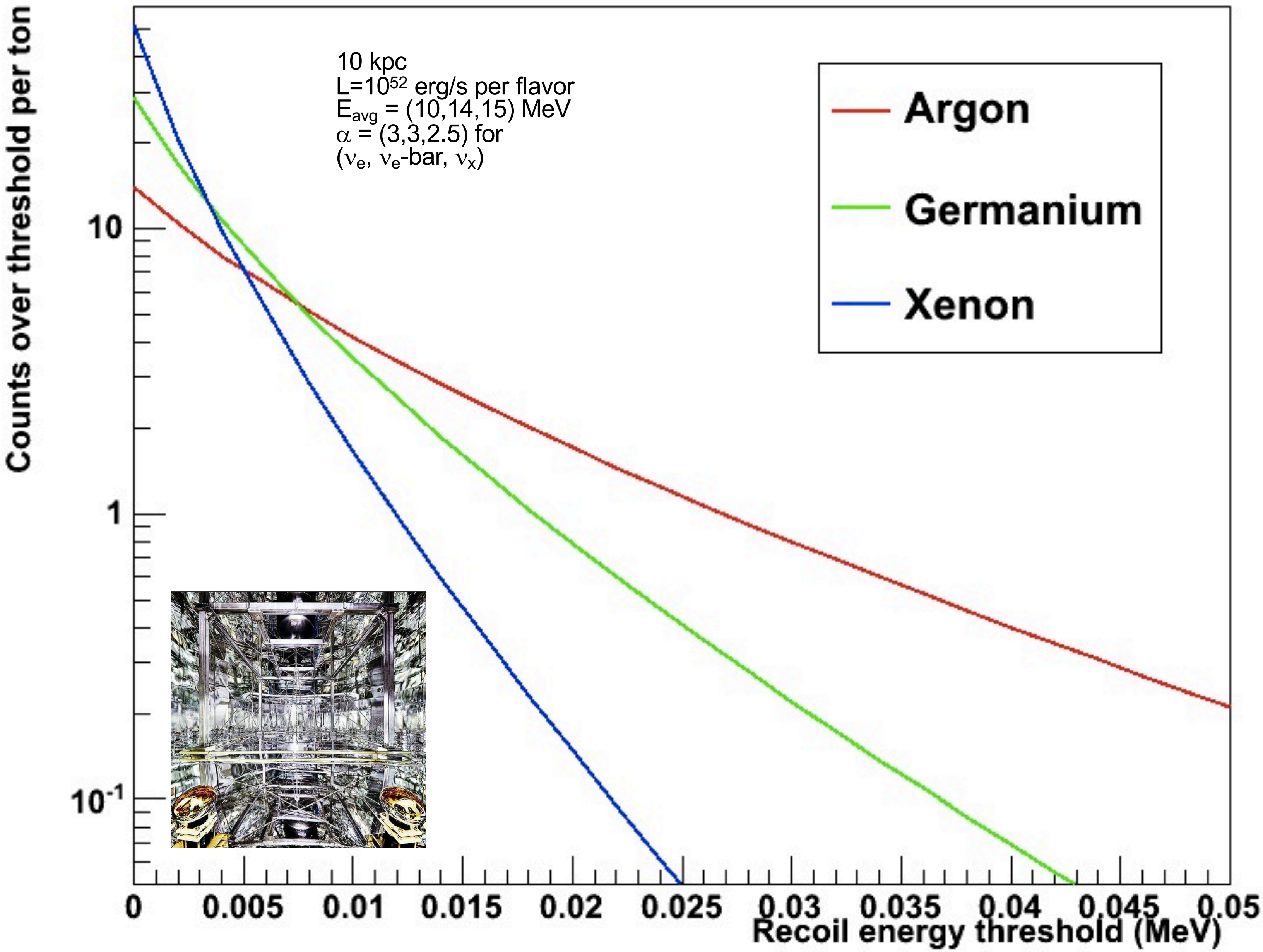
A. Ar, Ge, Xe

B. Xe, Ge, Ar

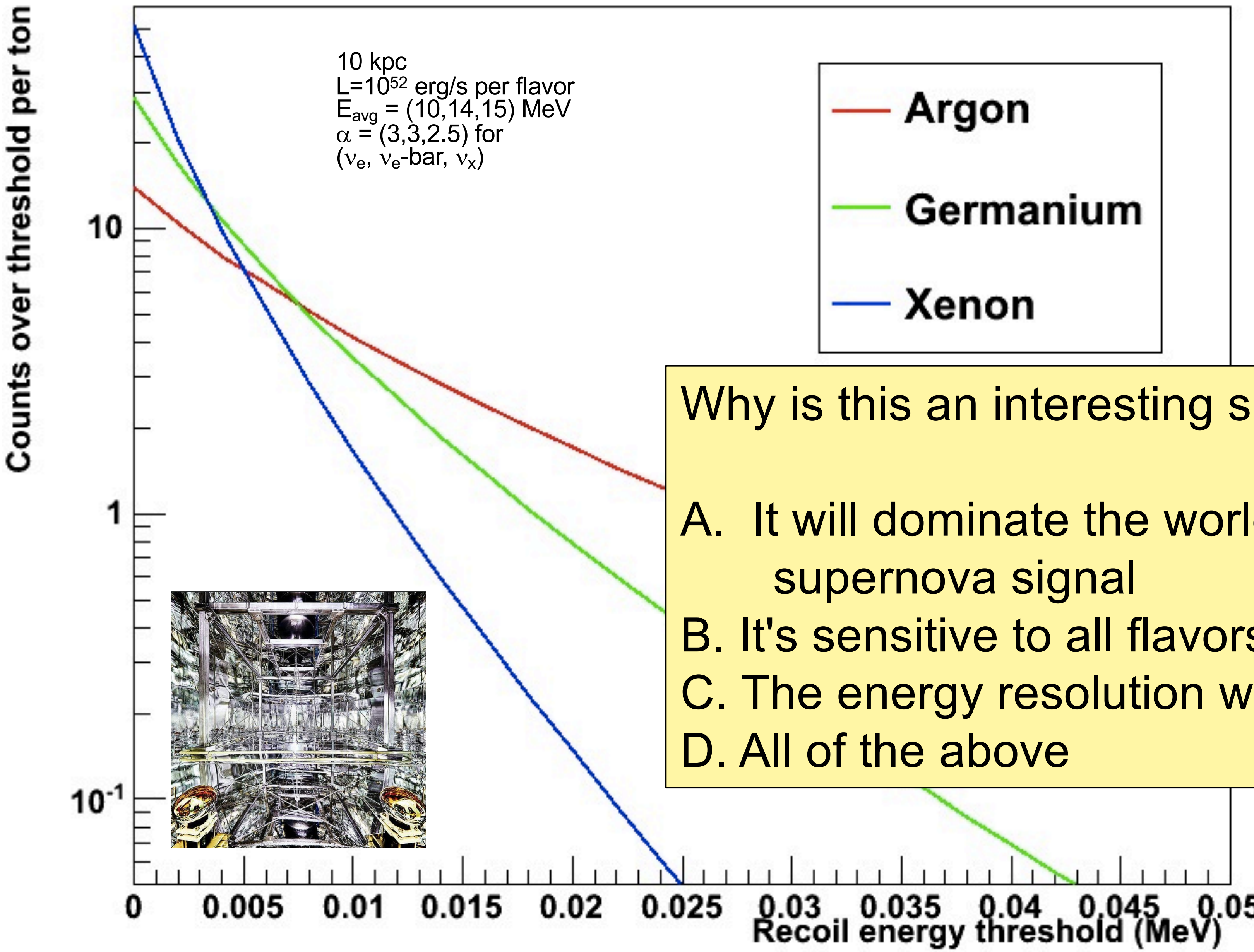
C. Ge, Xe, Ar

D. Ge, Ar, Xe

Supernova neutrinos in tonne-scale DM detectors



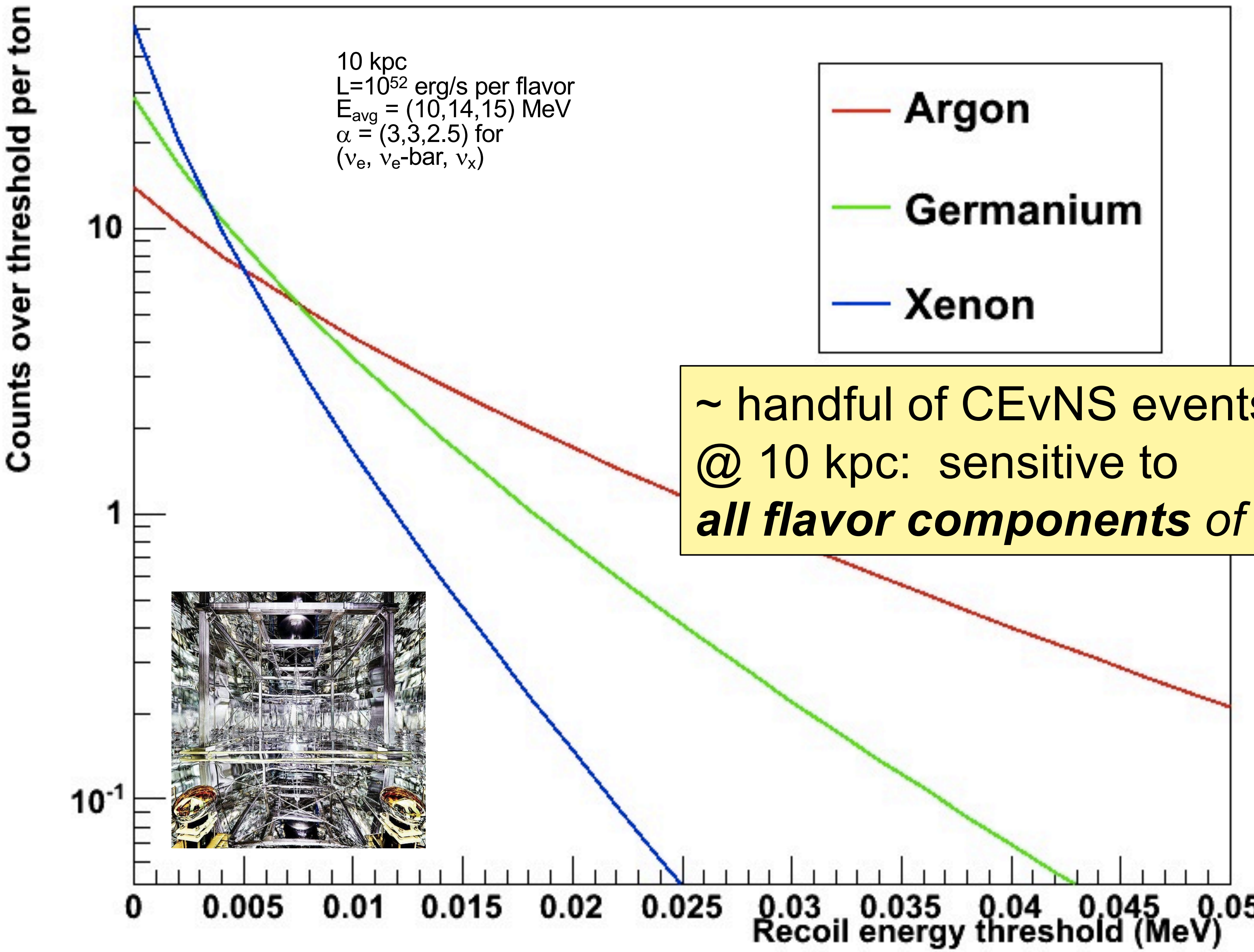
Supernova neutrinos in tonne-scale DM detectors



Why is this an interesting signal?

A. It will dominate the worldwide supernova signal
B. It's sensitive to all flavors of the burst
C. The energy resolution will be good
D. All of the above

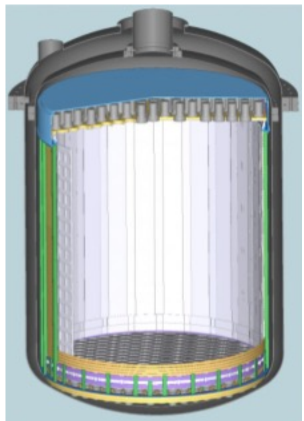
Supernova neutrinos in tonne-scale DM detectors

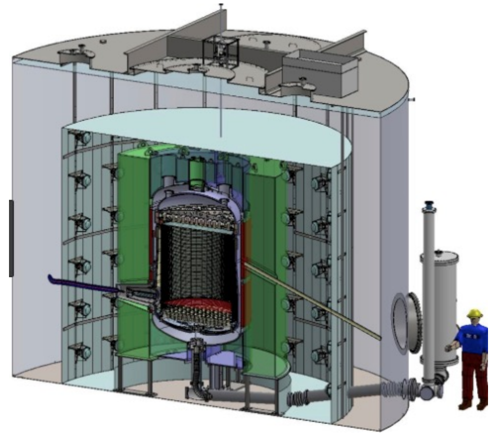


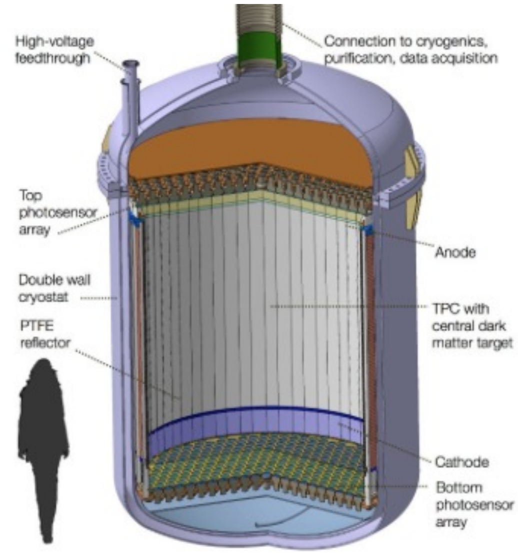
~ handful of CEvNS events per tonne @ 10 kpc: sensitive to ***all flavor components*** *of the flux*

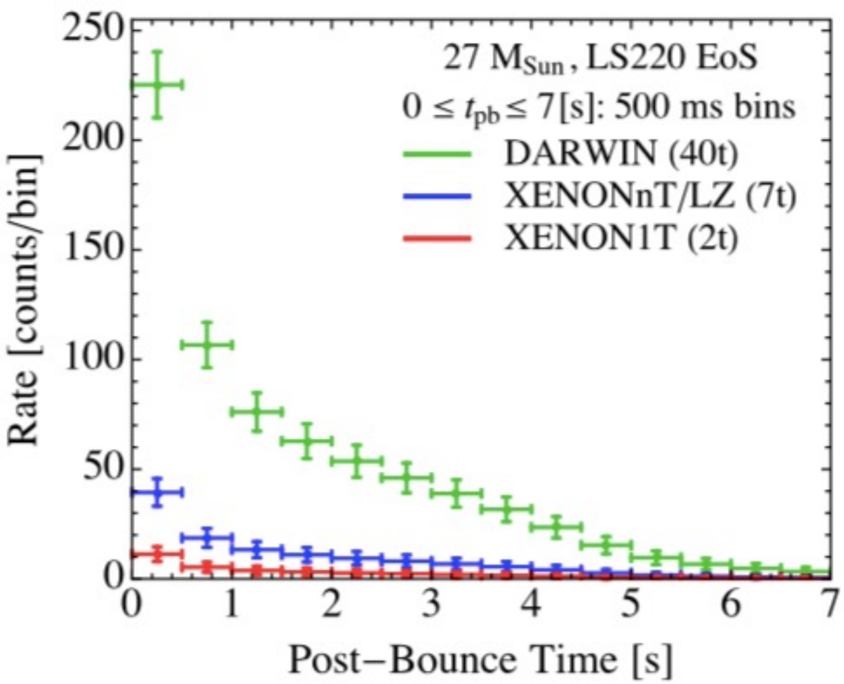
Detector example: **XENON/LZ/DARWIN**

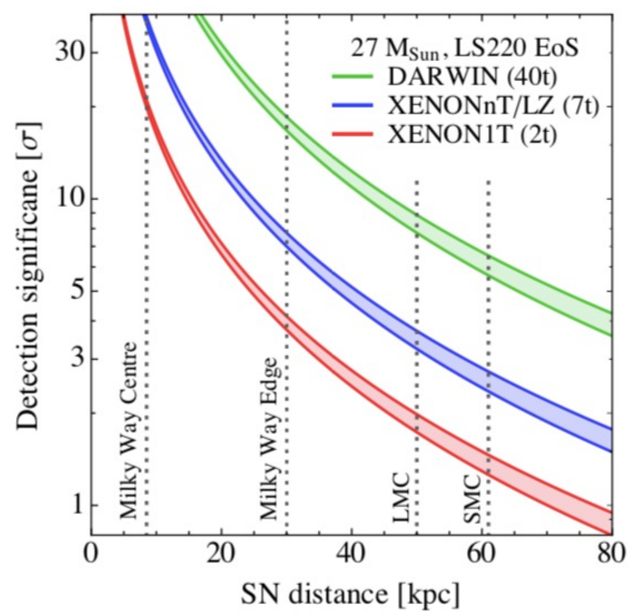
- dual-phase xenon time projection chambers









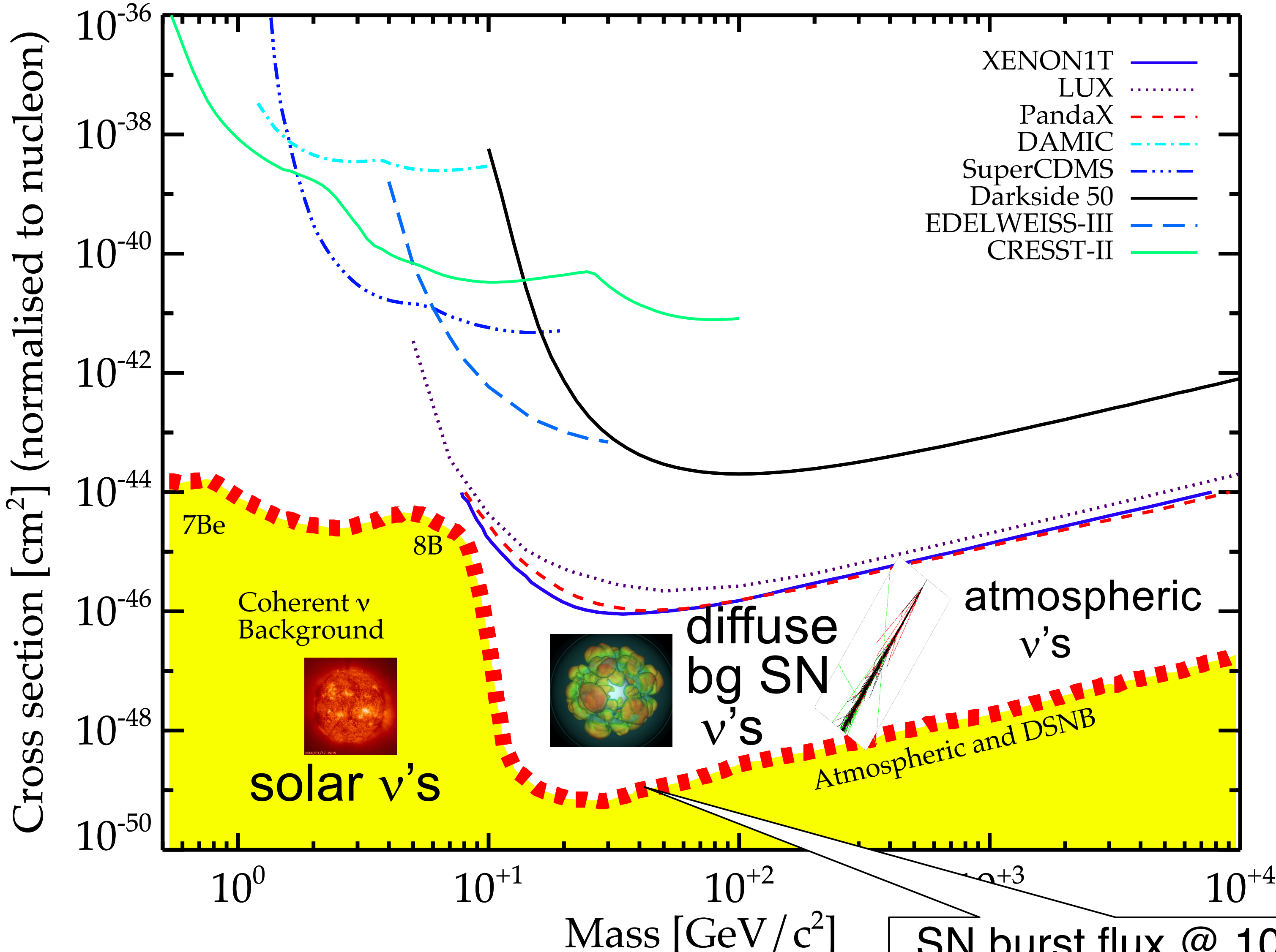


Lang et al.(2016). *Physical Review D*, *94*(10), 103009. http://doi.org/10.1103/PhysRevD.94.103009

The so-called “**neutrino floor**” for DM experiments

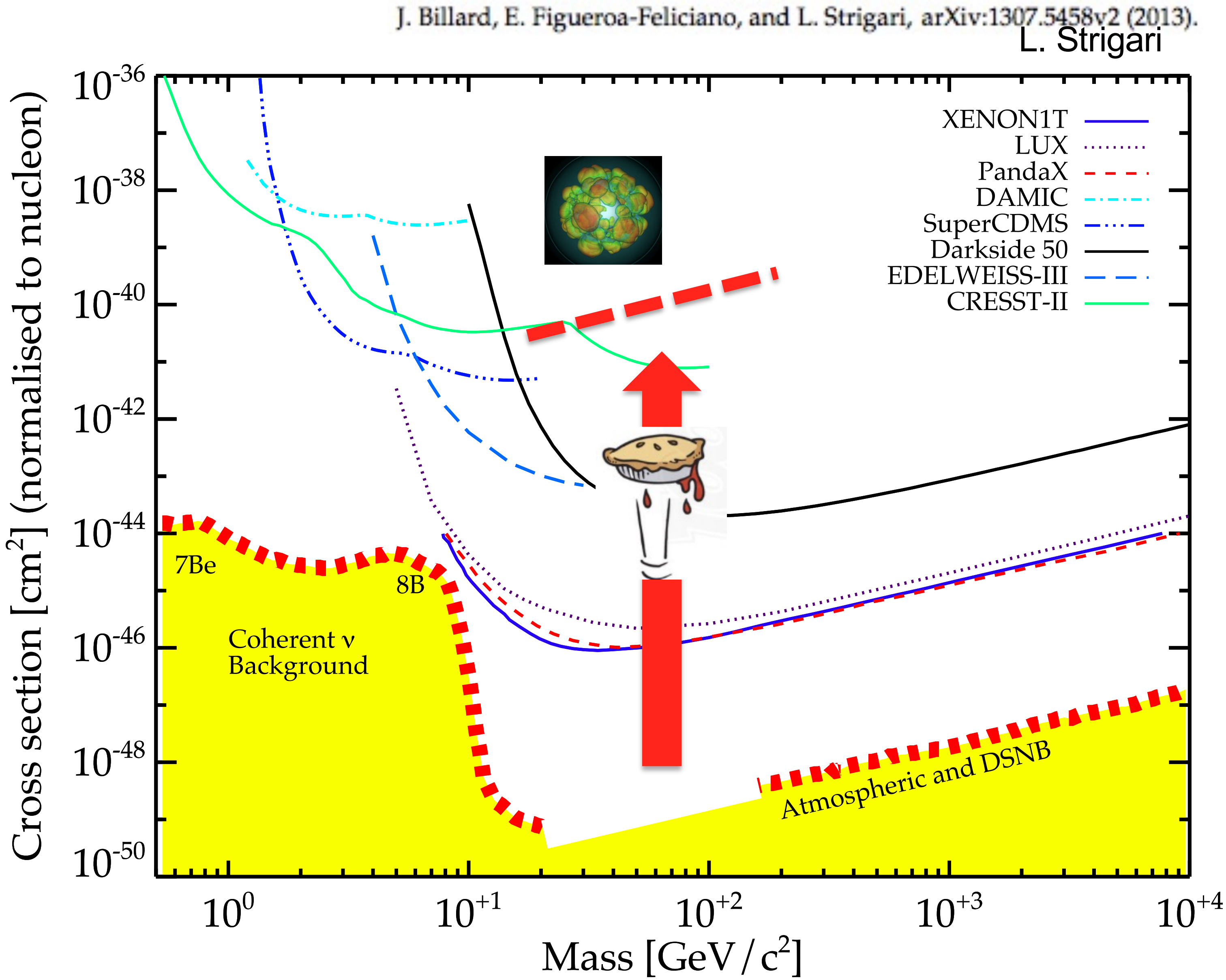
J. Billard, E. Figueroa-Feliciano, and L. Strigari, arXiv:1307.5458v2 (2013).

L. Strigari



SN burst flux @ 10 kpc is 9-10 orders of magnitude greater than DSNB flux

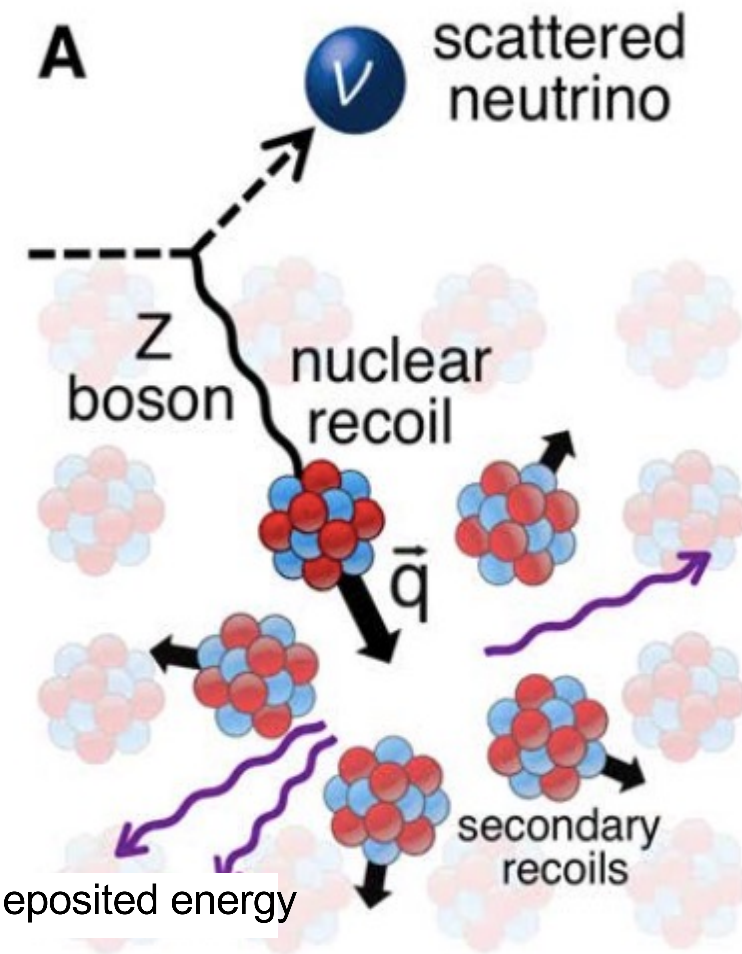
Think of a SN burst as “**the ν floor coming up to meet you**”



How to measure CEvNS

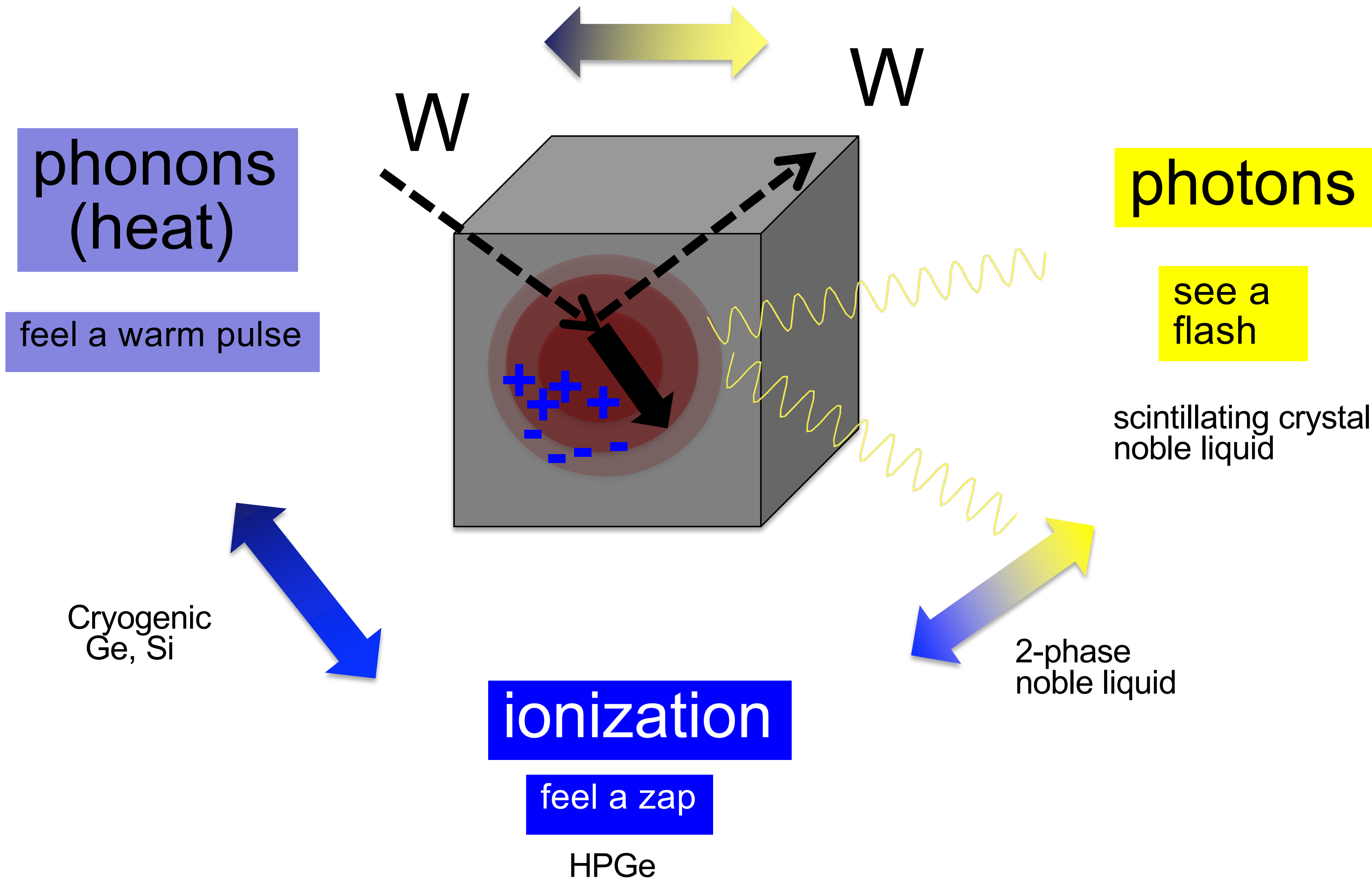
The only experimental signature:

tiny energy deposited by nuclear recoils in the target material



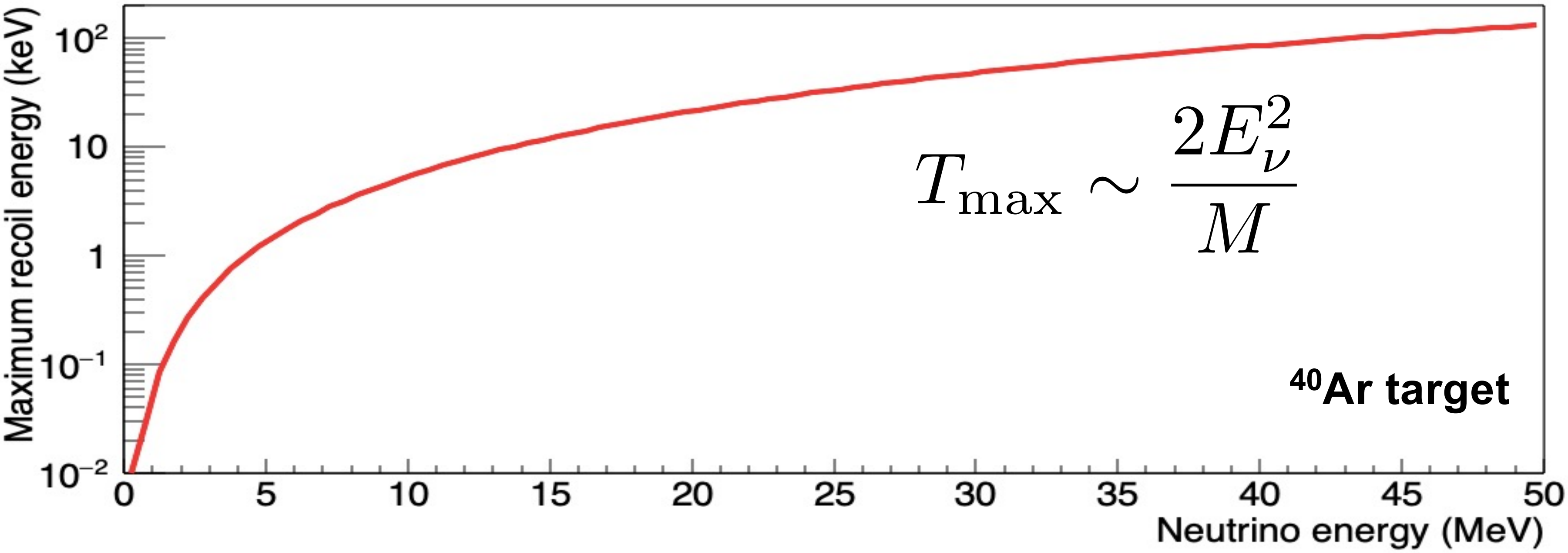
➔ detectors developed over the last ~few decades are sensitive to ~ keV to 10’s of keV recoils

Low-energy nuclear recoil detection strategies



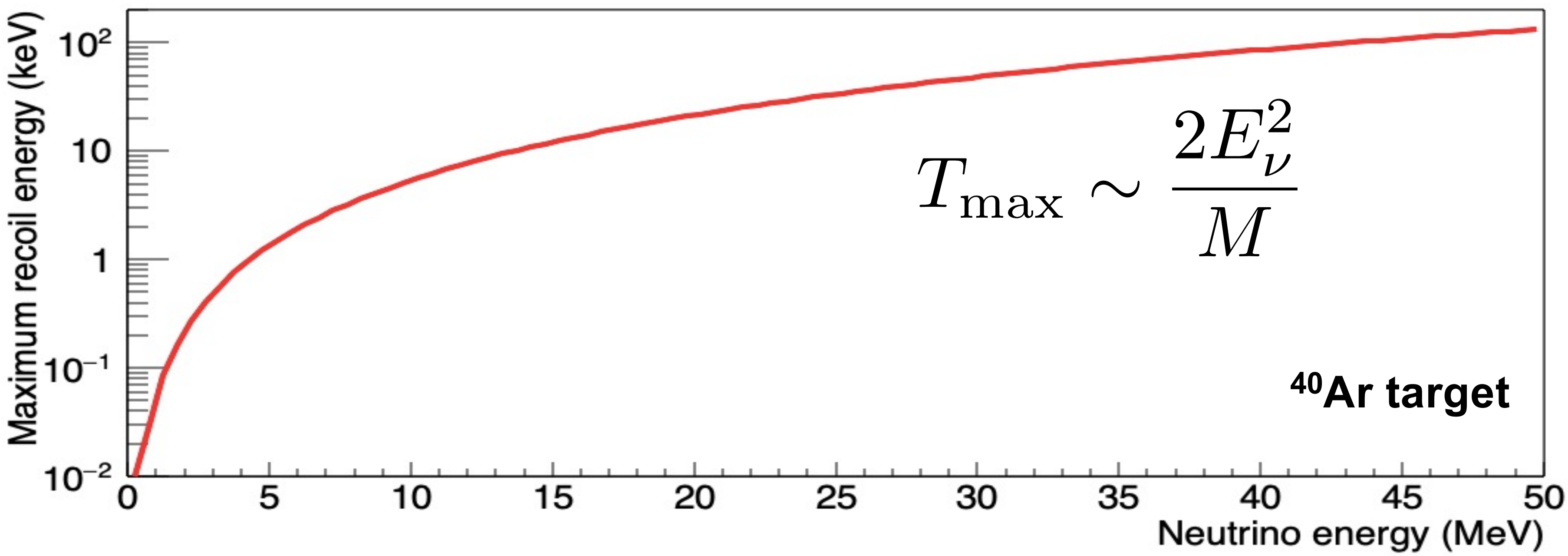
http://dmrc.snu.ac.kr/english/intro/intro1.html

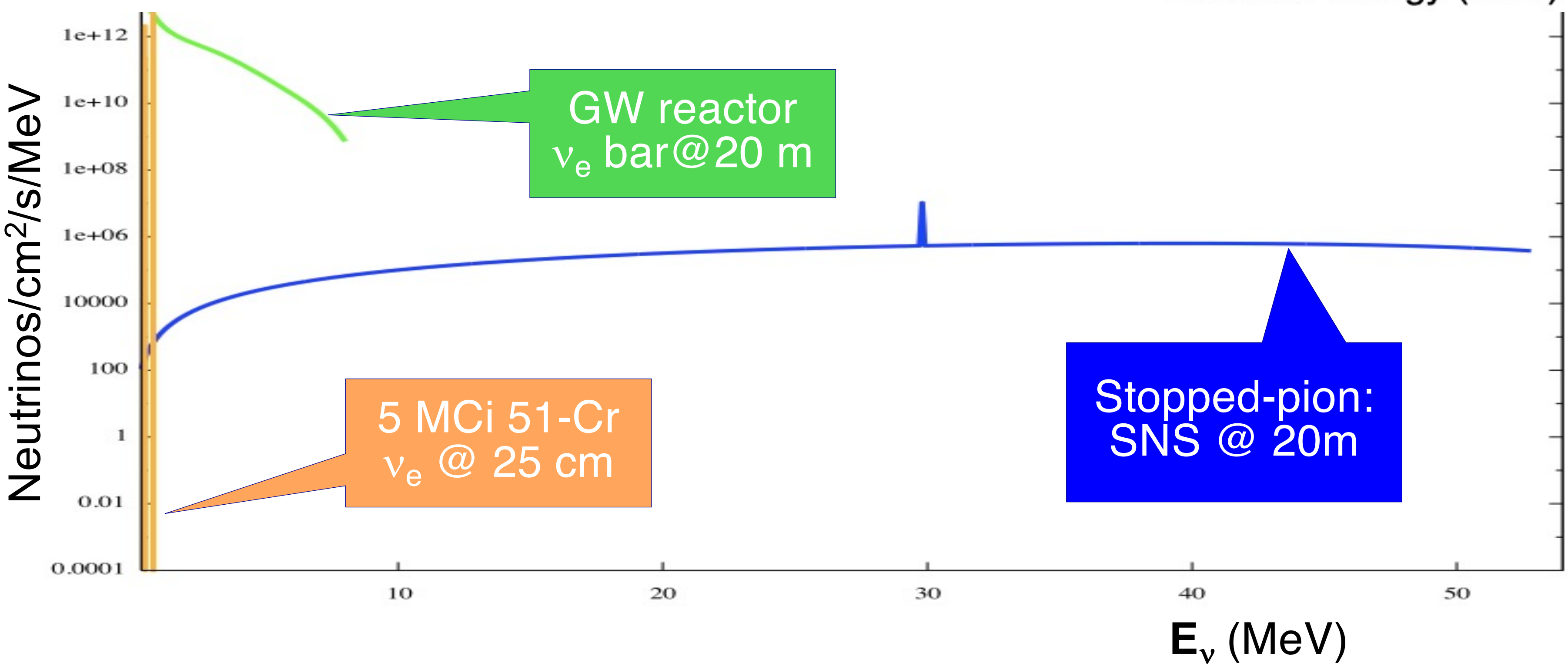
Maximum recoil energy as a function of E_ν



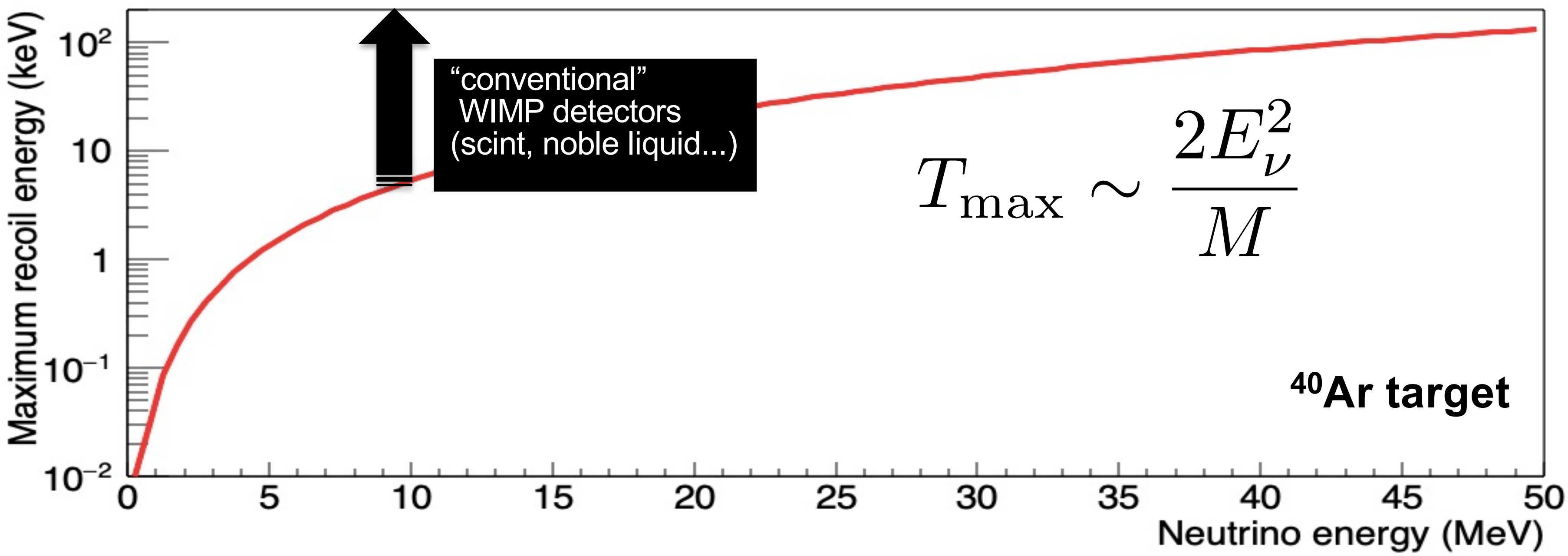
$$T_{\max} \sim \frac{2E_\nu^2}{M}$$

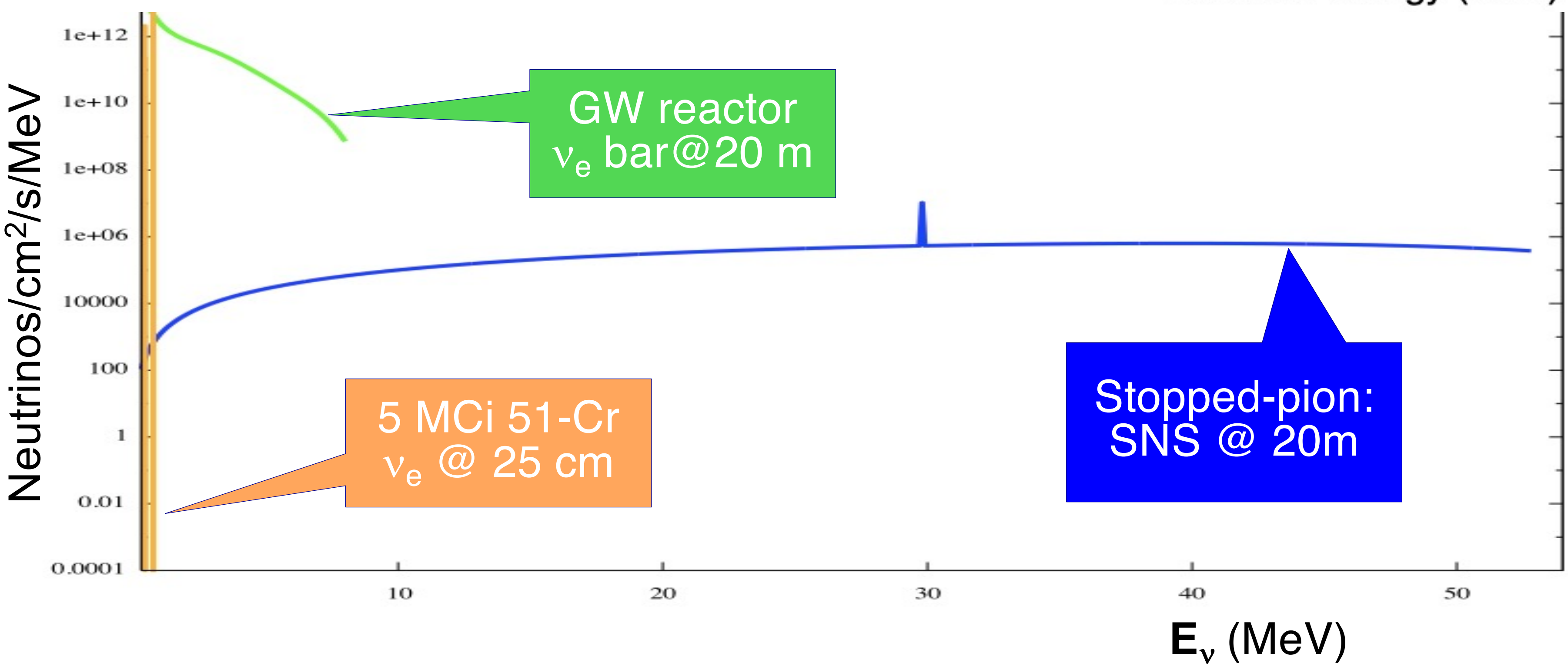
Maximum recoil energy as a function of E_ν



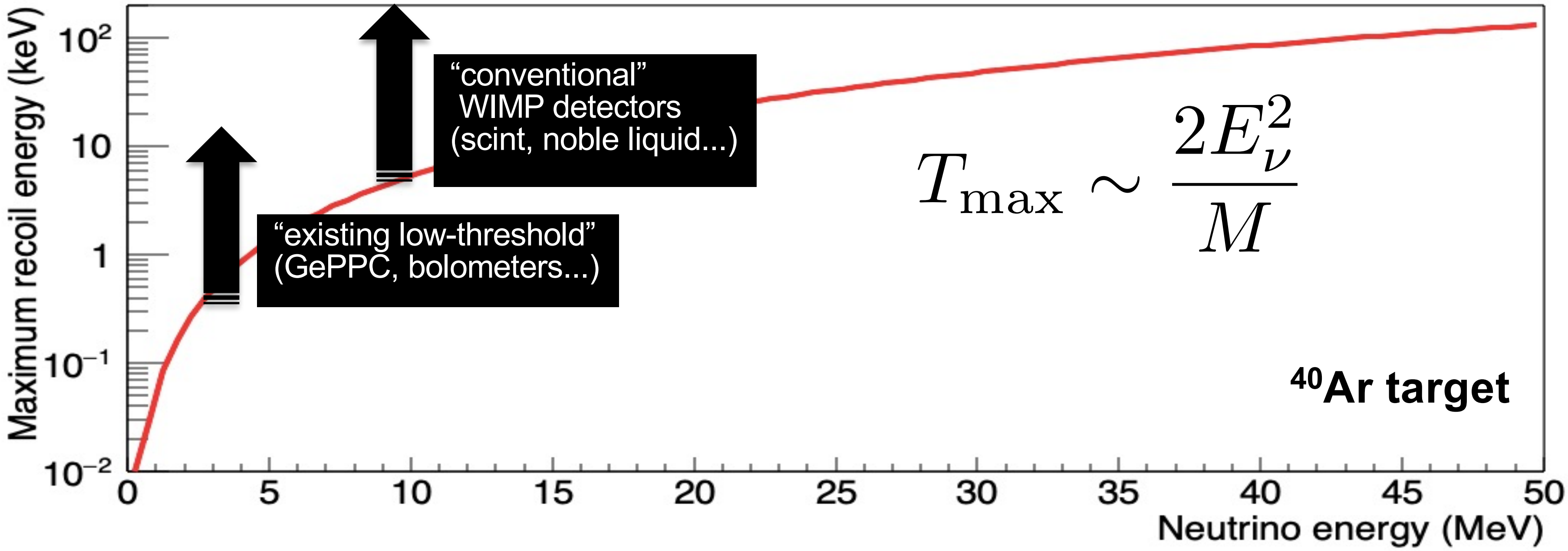


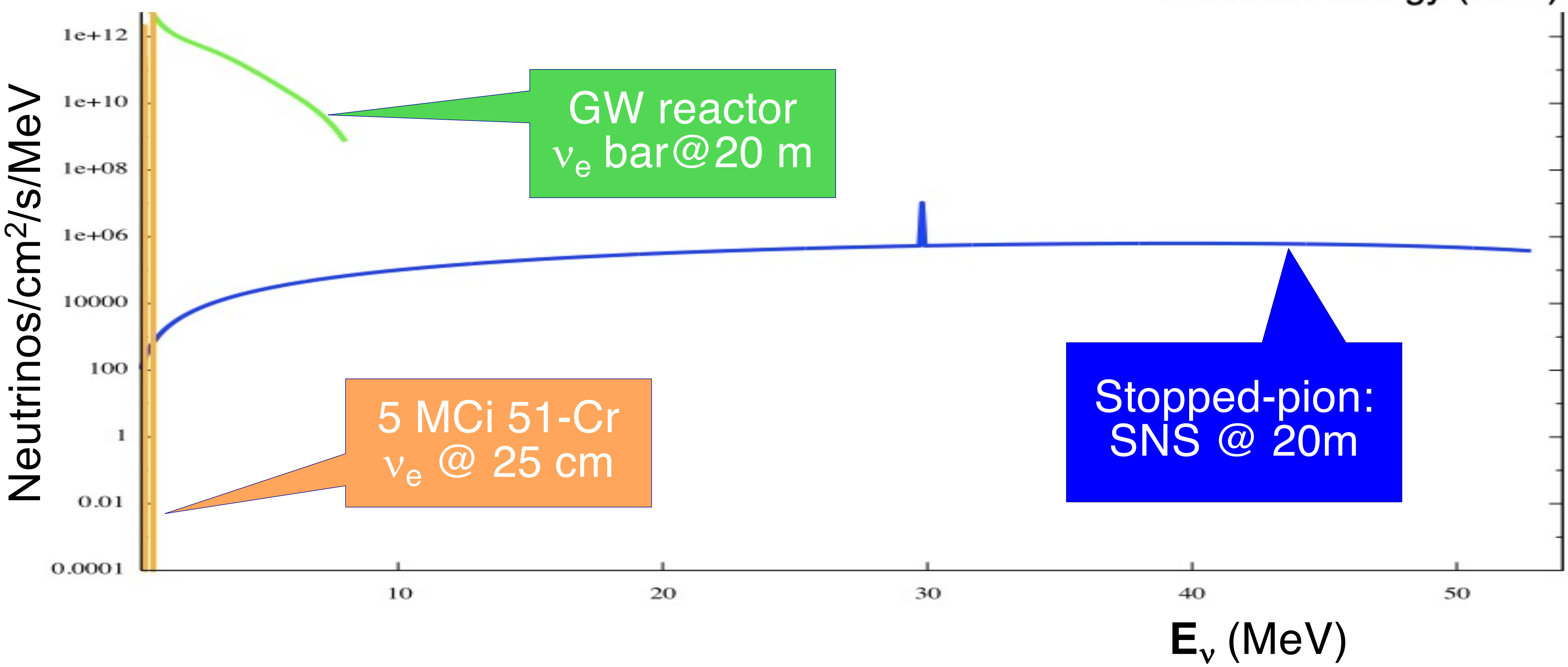
Maximum recoil energy as a function of E_ν



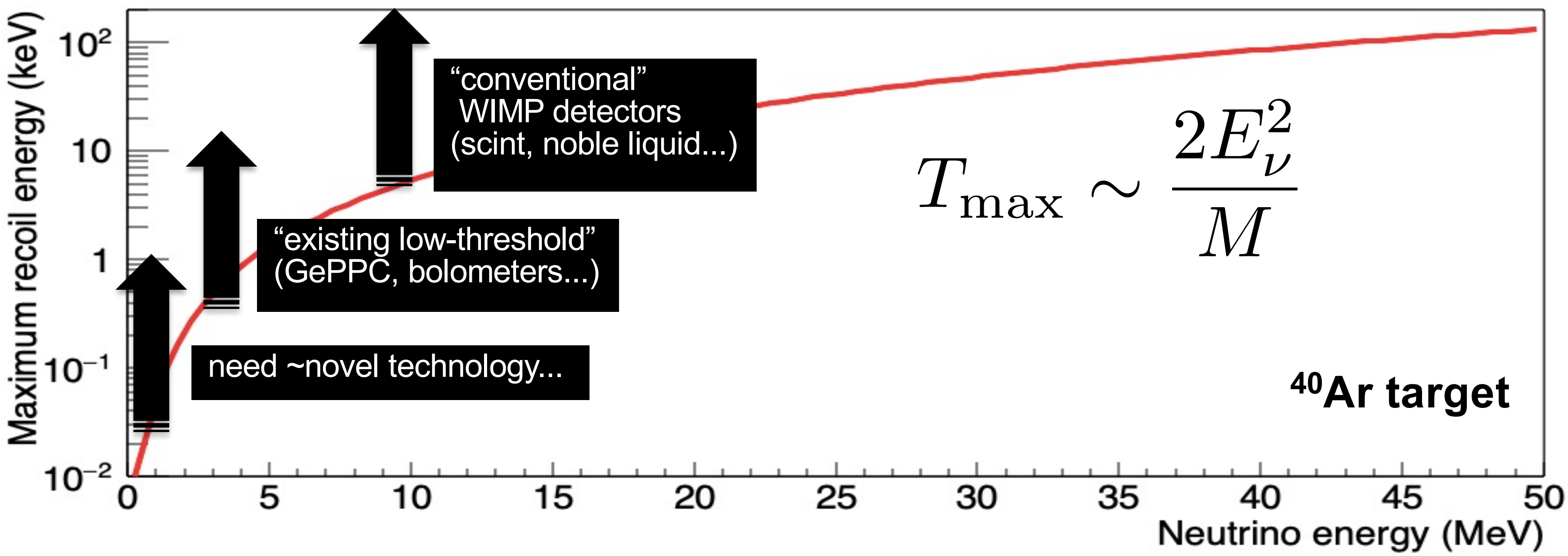


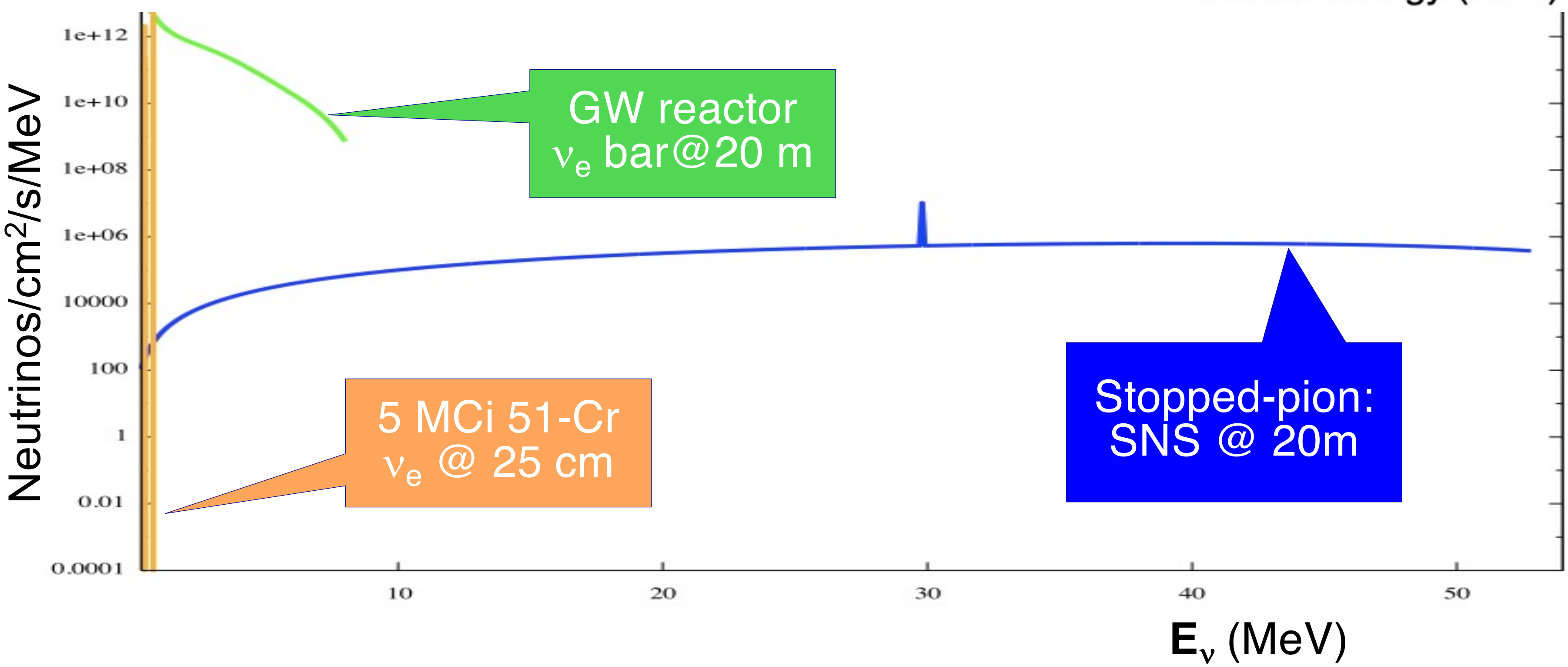
Maximum recoil energy as a function of E_ν



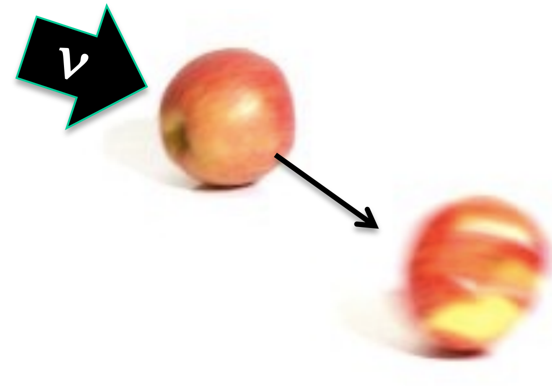


Maximum recoil energy as a function of E_ν





How to detect CEvNS?

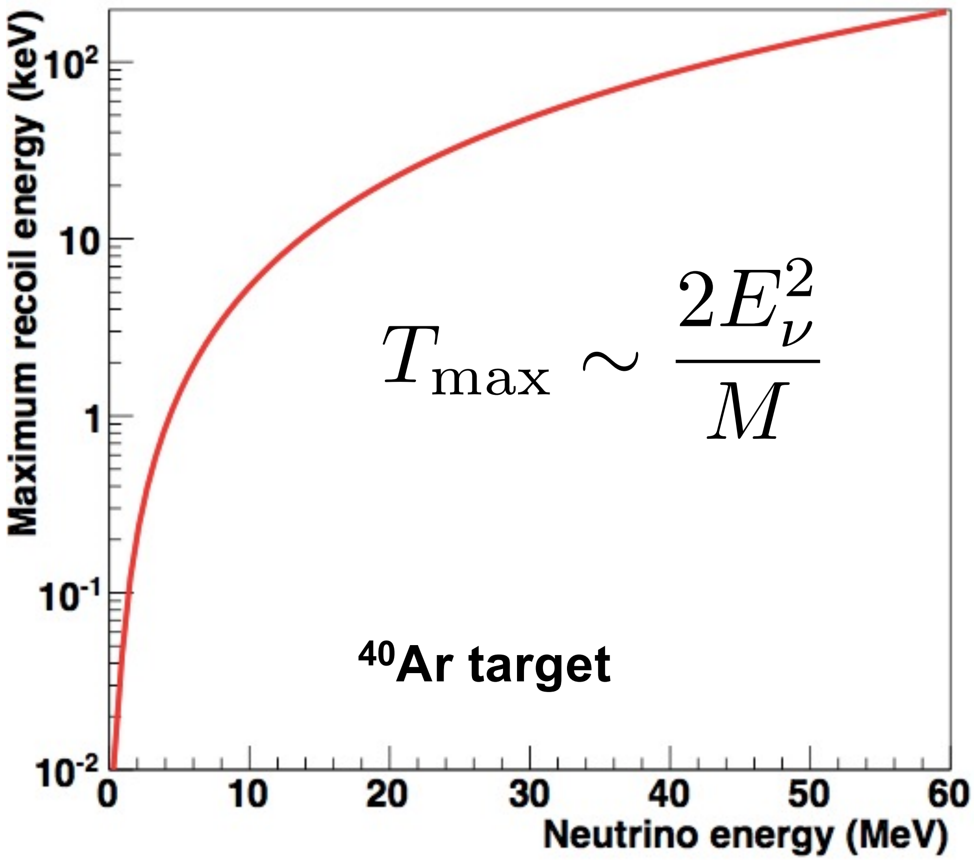


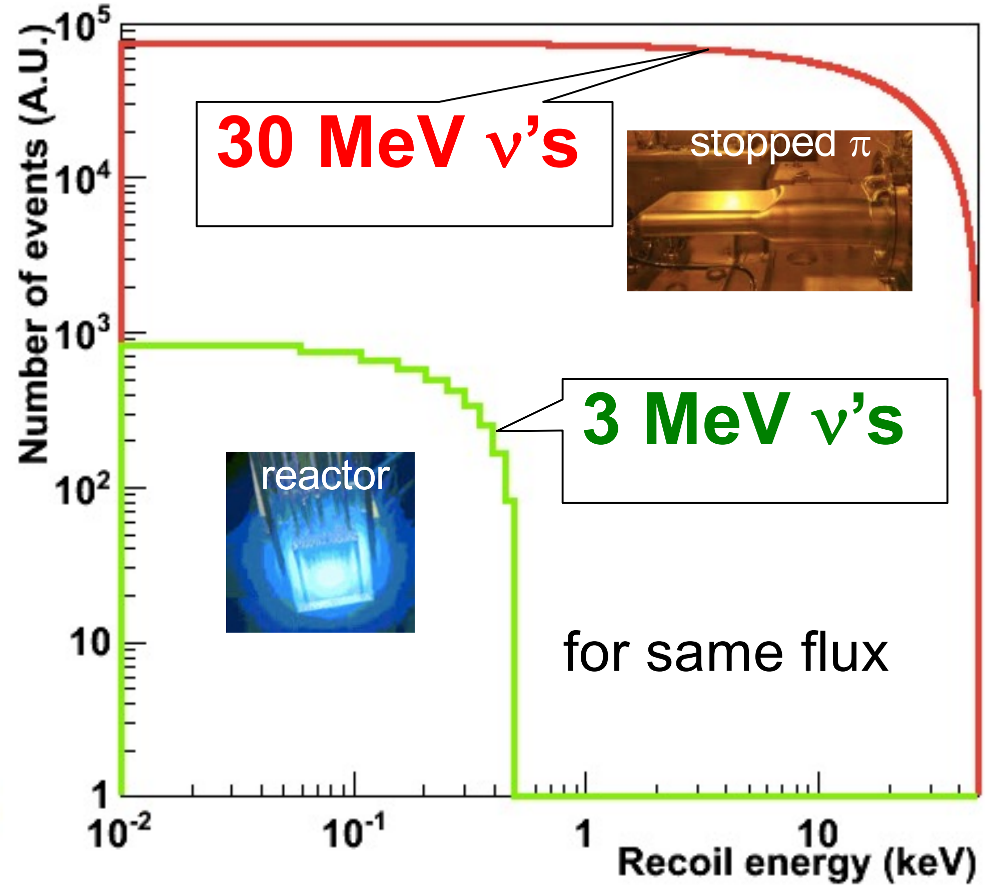
You need a neutrino source
and a detector

What do you want for your ν source?

- ✓ High flux
- ✓ Well understood spectrum
- ✓ Multiple flavors (physics sensitivity)
- ✓ Pulsed source if possible, for background rejection
- ✓ Ability to get close
- ✓ Practical things: access, control, ...

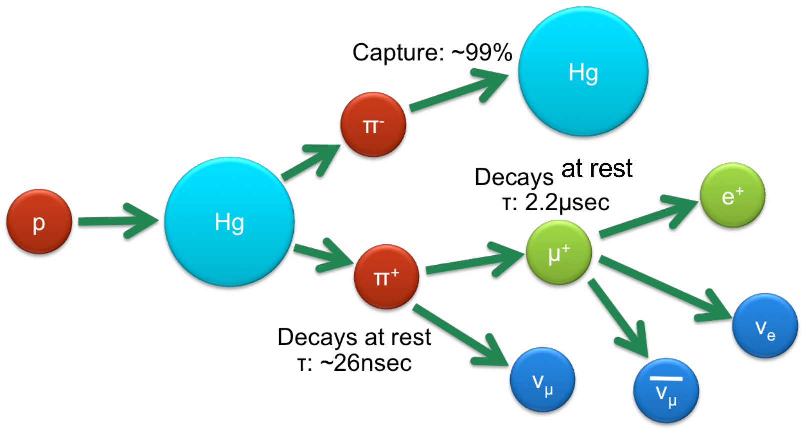
Both **cross-section** and **maximum recoil energy** **increase with neutrino energy**:

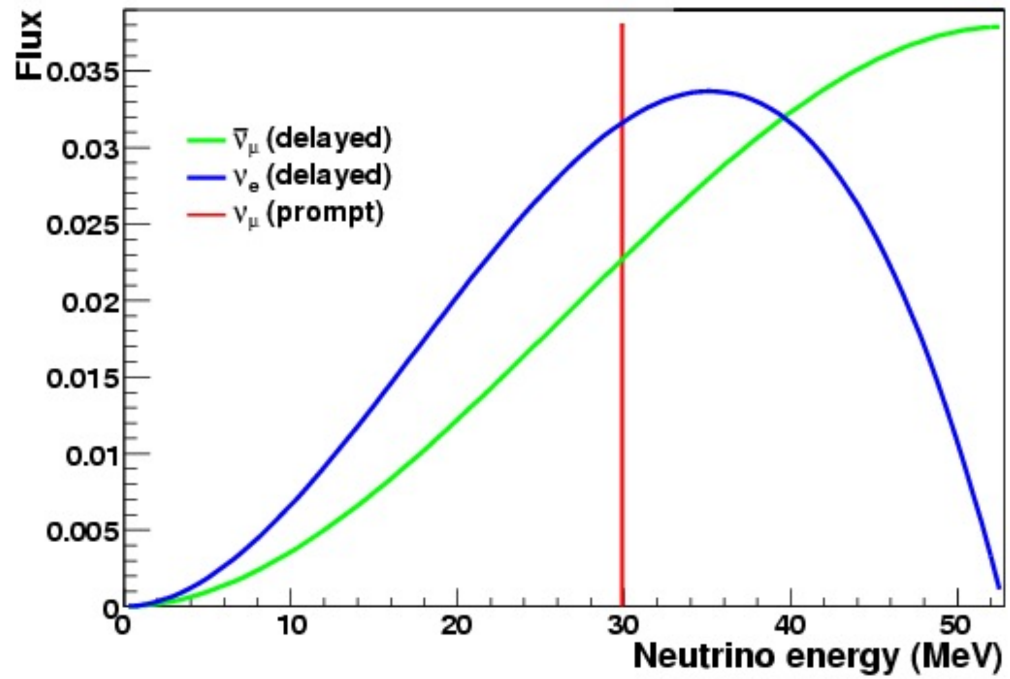




Want energy as large as possible while satisfying coherence condition: $Q \lesssim \frac{1}{R}$ (<~ 50 MeV for medium A)

Stopped-Pion (πDAR) Neutrinos





$$\pi^+ \rightarrow \mu^+ + \nu_\mu$$

2-body decay: monochromatic 29.9 MeV ν_μ
PROMPT

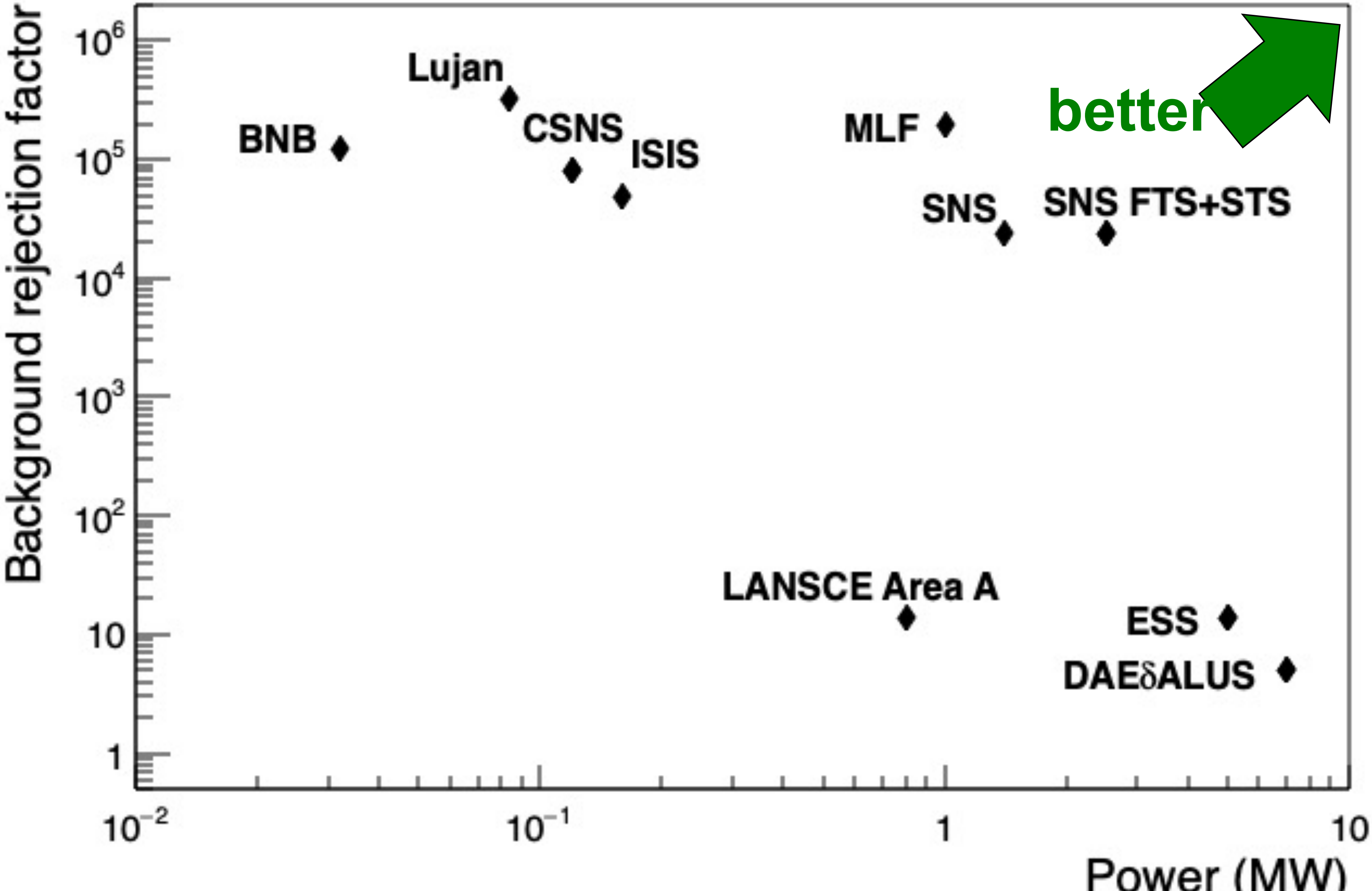
$$\downarrow$$

$$\mu^+ \rightarrow e^+ + \bar{\nu}_\mu + \nu_e$$

3-body decay: range of energies between 0 and $m_\mu/2$
DELAYED (2.2 μs)

from duty cycle

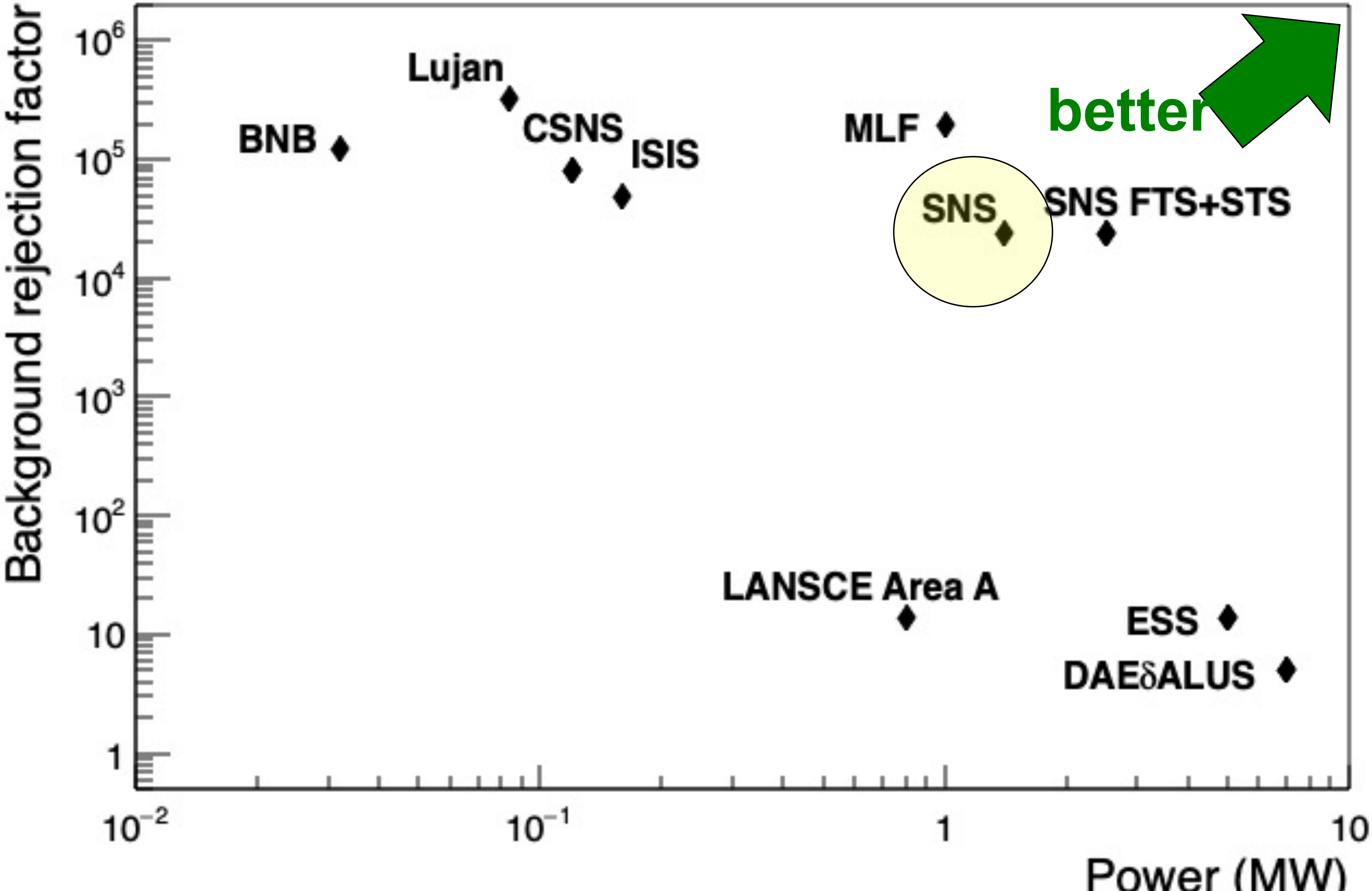
Comparison of pion decay-at-rest ν sources



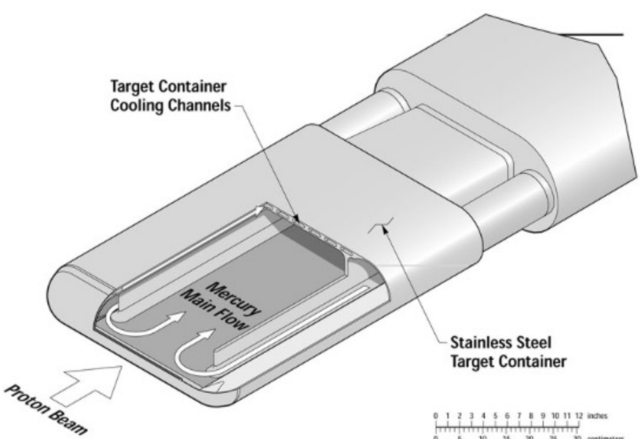
$\propto \nu$ flux

Comparison of pion decay-at-rest ν sources

from duty cycle



$\propto \nu$ flux

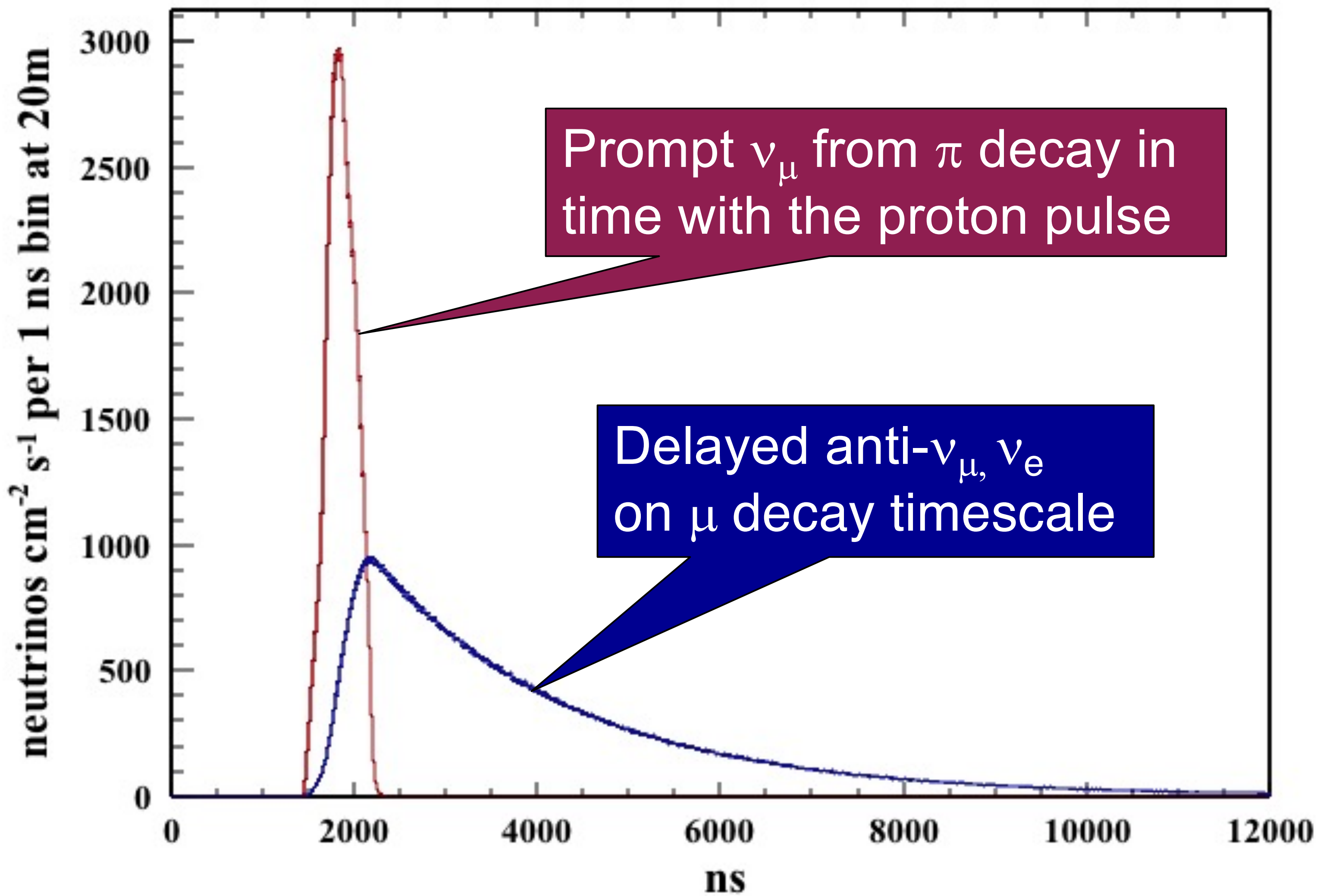


Proton beam energy: 0.9-1.3 GeV
Total power: 0.9-1.4 MW
Pulse duration: 380 ns FWHM
Repetition rate: 60 Hz
Liquid mercury target

The neutrinos are free!

Time structure of the SNS source

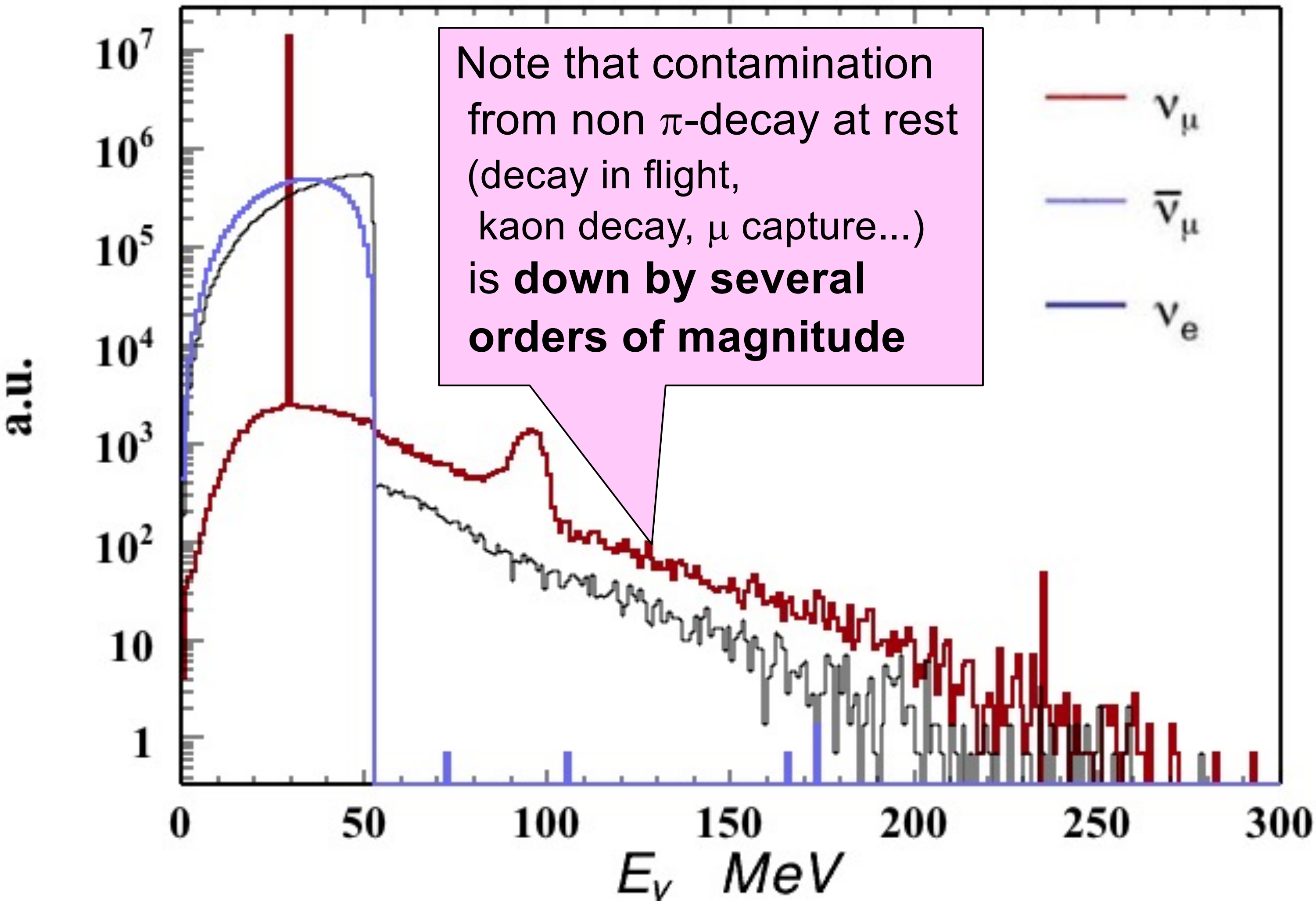
60 Hz *pulsed* source



Background rejection factor ~few x 10^{-4}

The SNS has **large, extremely clean** stopped-pion ν flux

0.08 neutrinos per flavor per proton on target



SNS flux (1.4 MW):
430 x 10^5 ν/cm²/s
@ 20 m

The COHERENT collaboration

http://sites.duke.edu/coherent

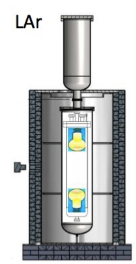
~90 members,
20 institutions
4 countries

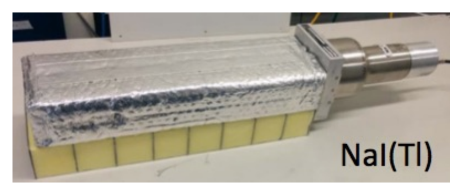
arXiv:1509.08702

COHERENT CEvNS Detectors

Nuclear Target	Technology		Mass (kg)	Distance from source (m)	Recoil threshold (keVr)
CsI[Na]	Scintillating crystal	flash	14.6	19.3	6.5
Ge	HPGe PPC	zap	18	22	<few
LAr	Single-phase	flash	24	27.5	20
NaI[Tl]	Scintillating crystal	flash	185*/3338	25	13

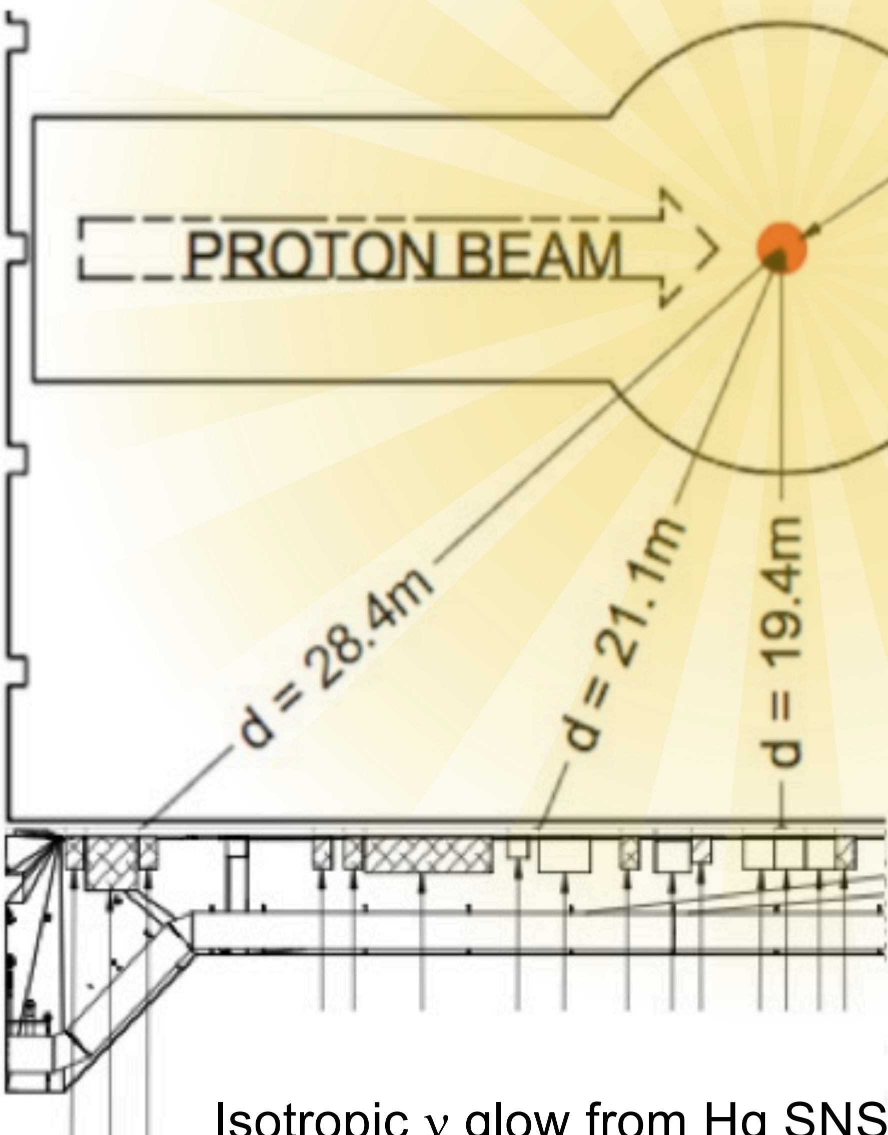
Multiple detectors for N^2 dependence of the cross section





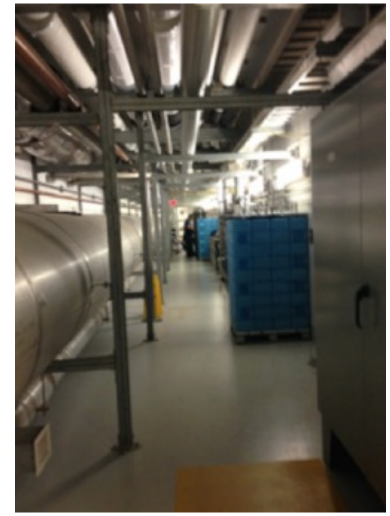
Siting for deployment in SNS basement

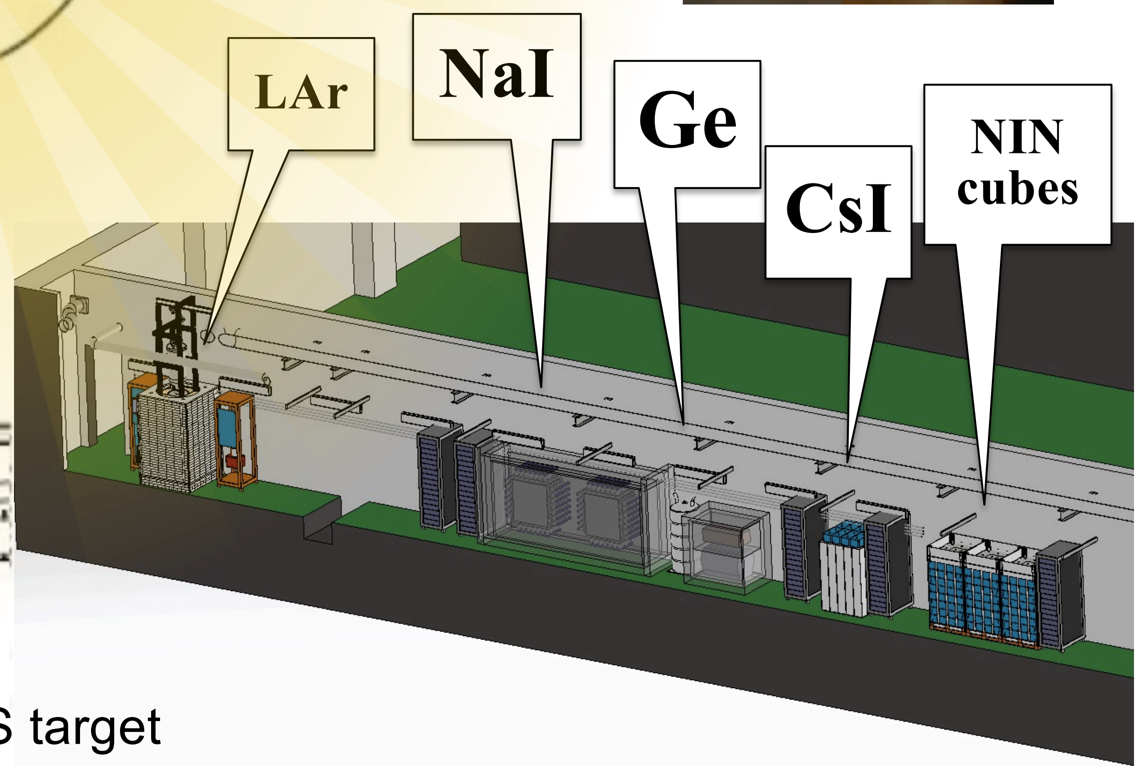
(measured neutron backgrounds low,
~ 8 mwe overburden)



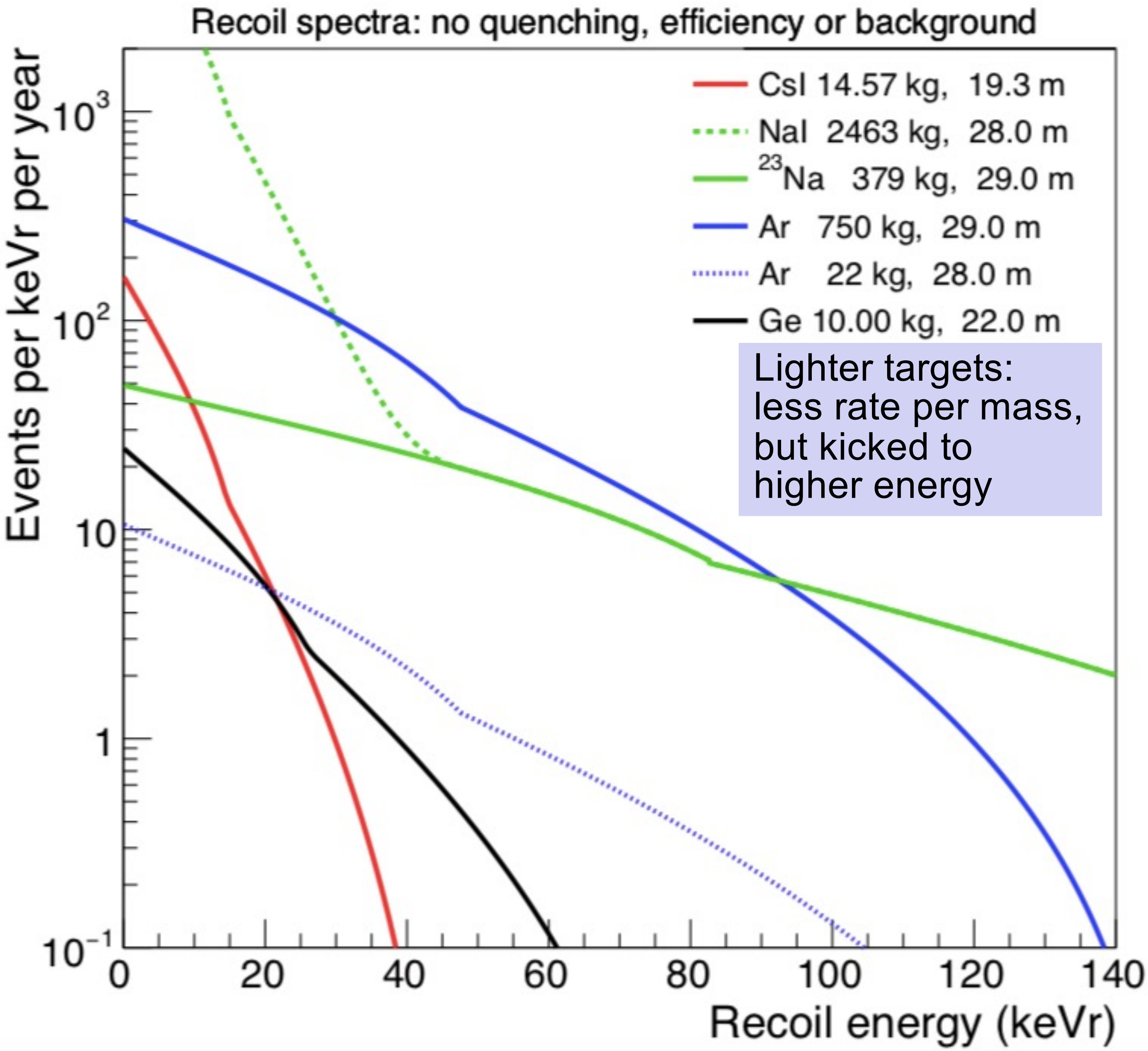
Isotropic ν glow from Hg SNS target

View looking
down "Neutrino Alley"





Expected recoil energy distribution



If 100 counts are expected in 10 kg of argon at 20 m, how many are expected in 100 kg at 40 m?

A. 2.5

B. 25

C. 100

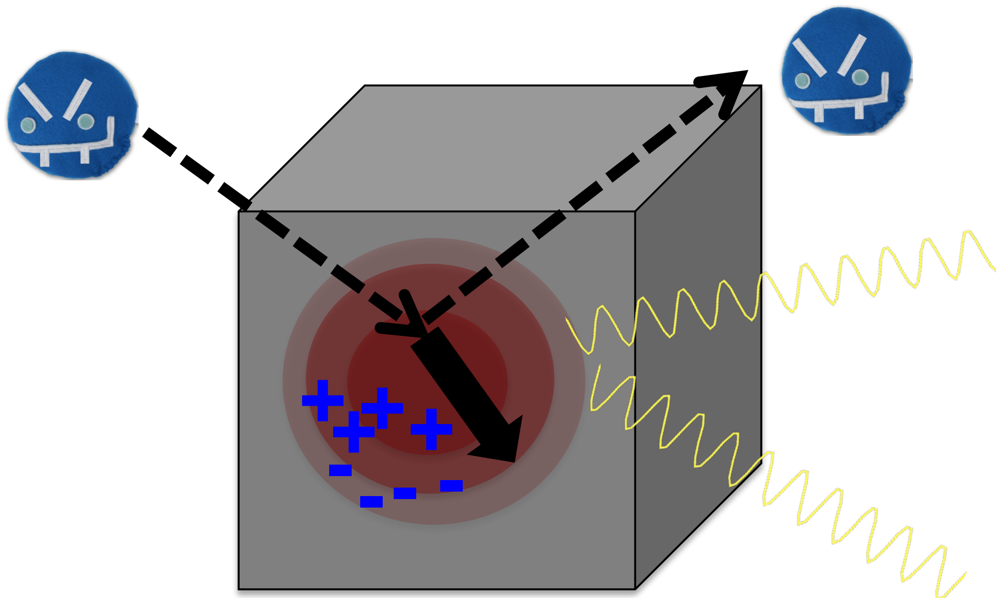
D. 250

E. 2500

Backgrounds

Usual suspects:
- cosmogenics
- ambient and intrinsic radioactivity
- detector-specific noise and dark rate

Neutrons are especially not your friends*



Steady-state backgrounds can be *measured* off-beam-pulse
... in-time backgrounds must be carefully characterized

*Thanks to Robert Cooper for the "mean neutron"

A “friendly fire” in-time background:

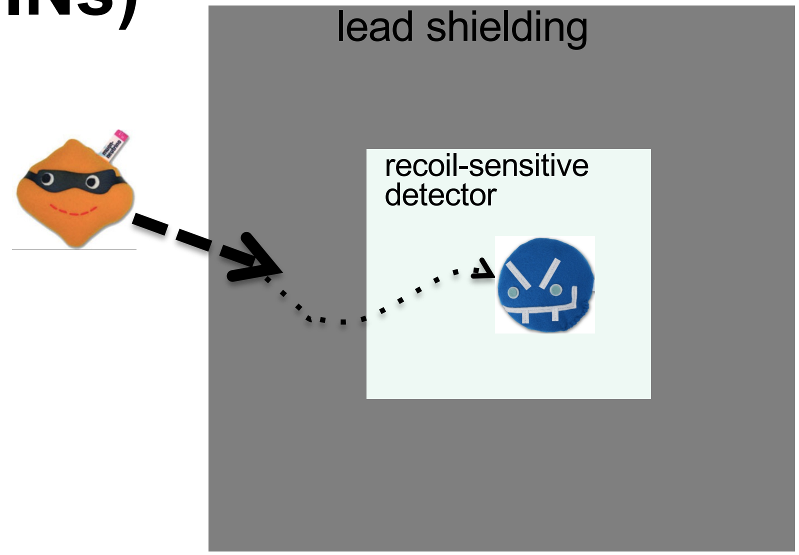
Neutrino Induced Neutrons (NINs)

$$\nu_e + {}^{208}\mathrm{Pb} \rightarrow {}^{208}\mathrm{Bi}^* + e^- \quad \mathrm{CC}$$

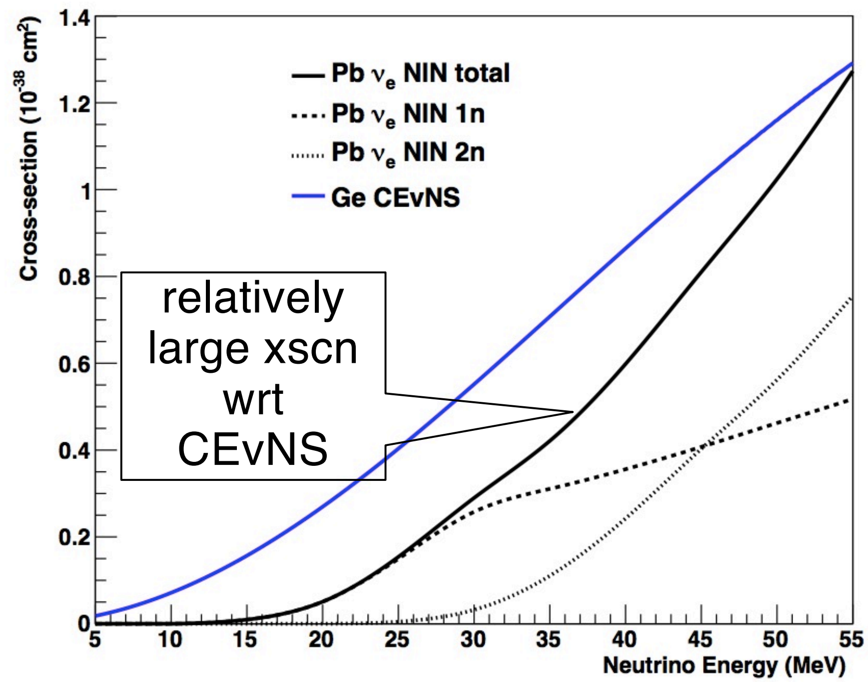
1n, 2n emission

$$\nu_x + {}^{208}\mathrm{Pb} \rightarrow {}^{208}\mathrm{Pb}^* + \nu_x \quad \mathrm{NC}$$

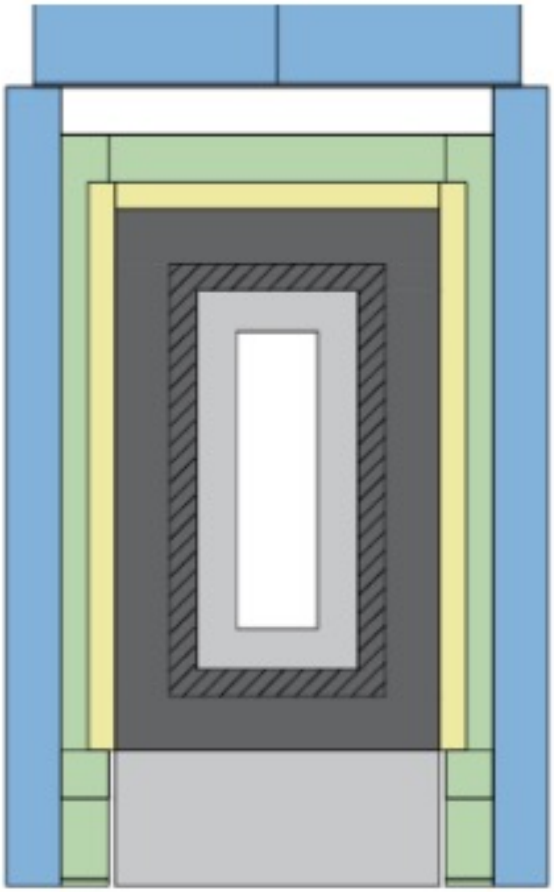
1n, 2n, γ emission

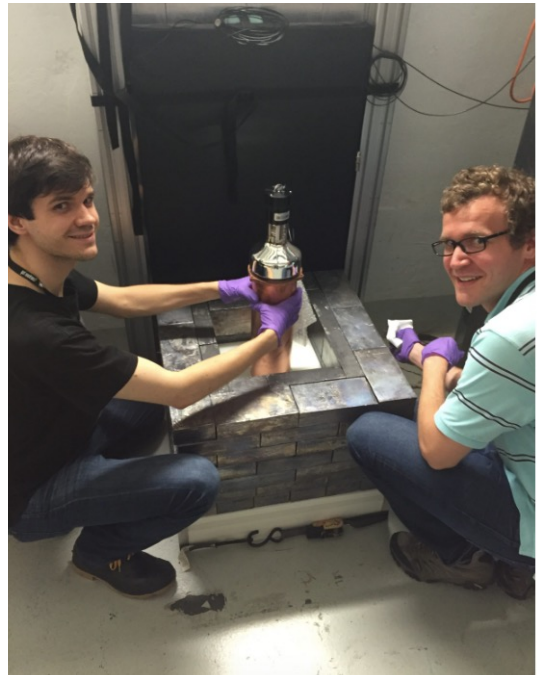


- potentially non-negligible background from shielding
- requires careful shielding design
- large uncertainties (factor of few) in xscn calculation
- [Also: a signal in itself, e.g, HALO SN detector]

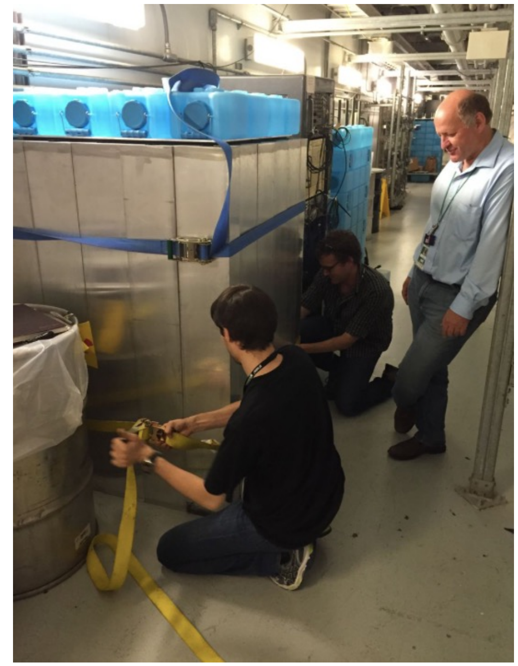


The CsI Detector in Shielding in Neutrino Alley at the SNS





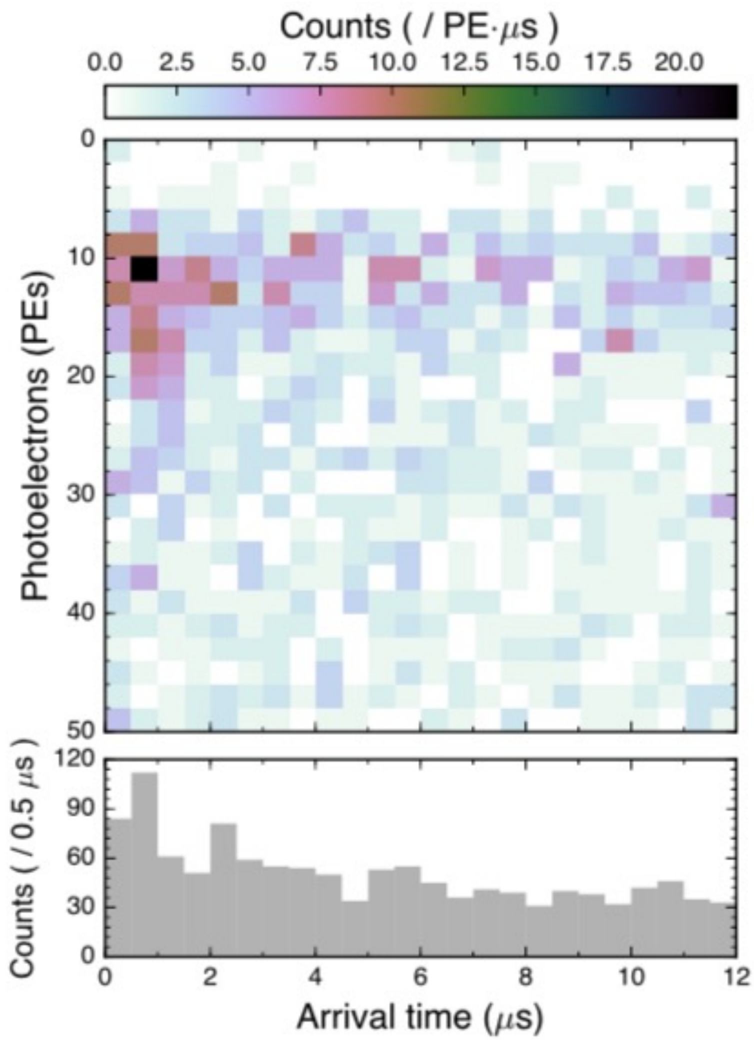
A hand-held detector!



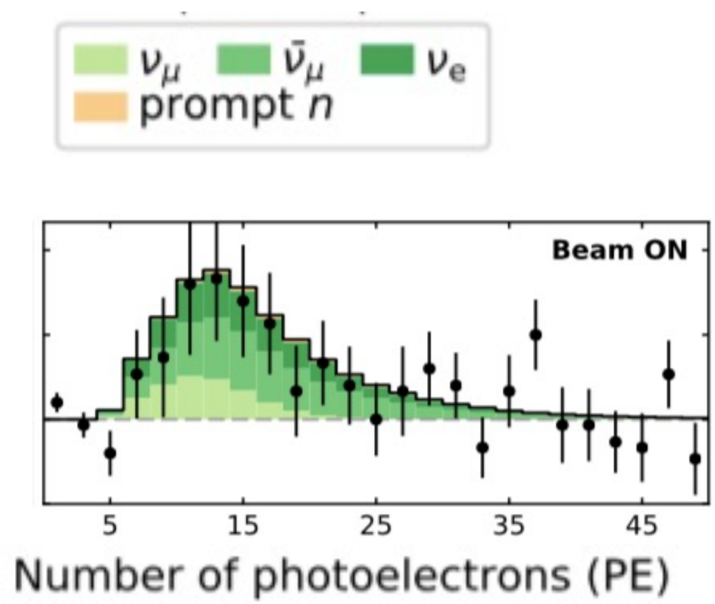
Almost wrapped up...

Layer	HDPE*	Low backg. lead	Lead	Muon veto	Water
Thickness	3"	2"	4"	2"	4"
Colour					

First light at the SNS (stopped-pion neutrinos) with 14.6-kg CsI[Na] detector



DOI: 10.5281/zenodo.1228631



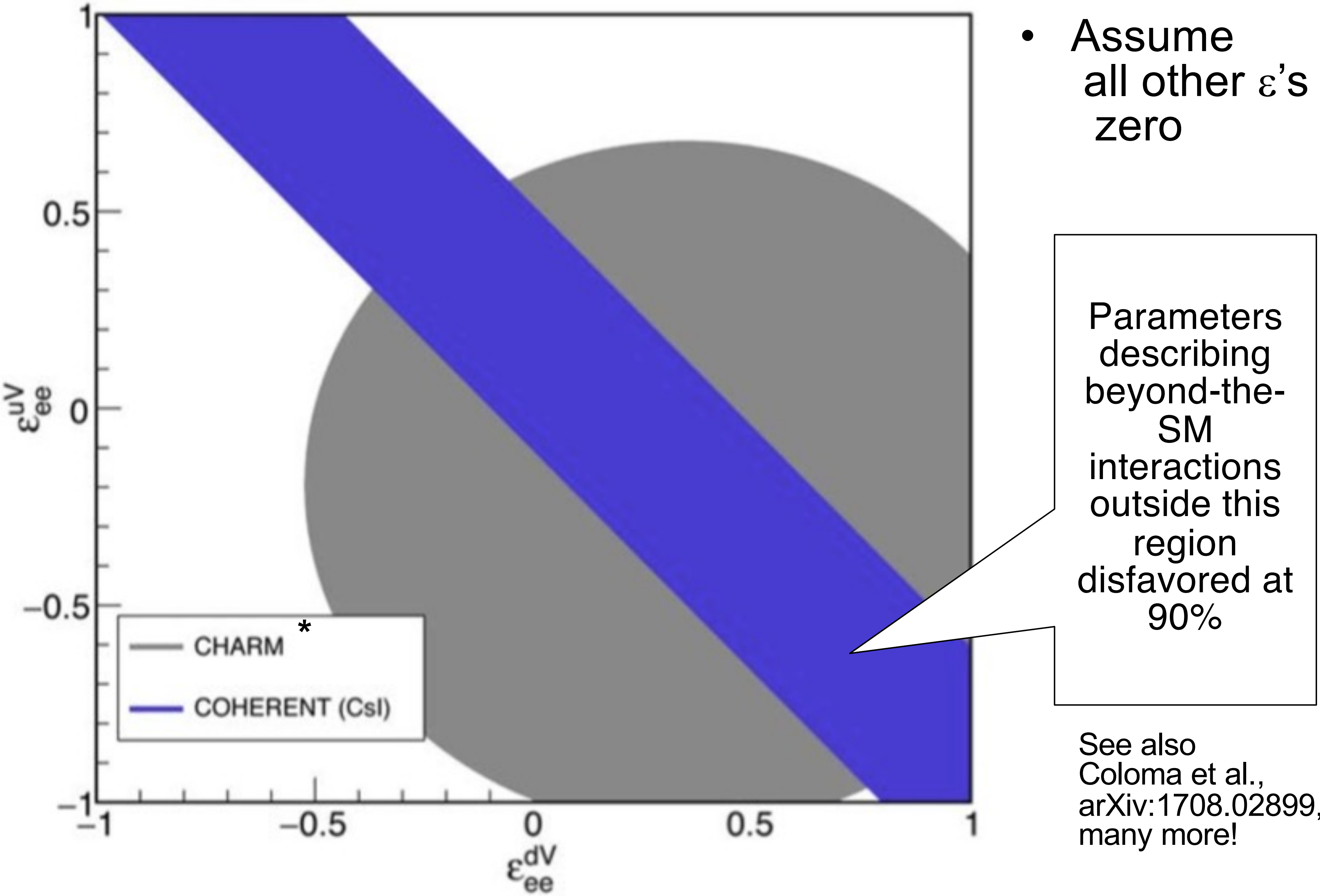
Background-subtracted and integrated over time

$$\mathrm{PE} \propto T \propto Q^2$$

→ measure of the Q spectrum

D. Akimov et al., *Science*, 2017
http://science.sciencemag.org/content/early/2017/08/02/science.aao0990

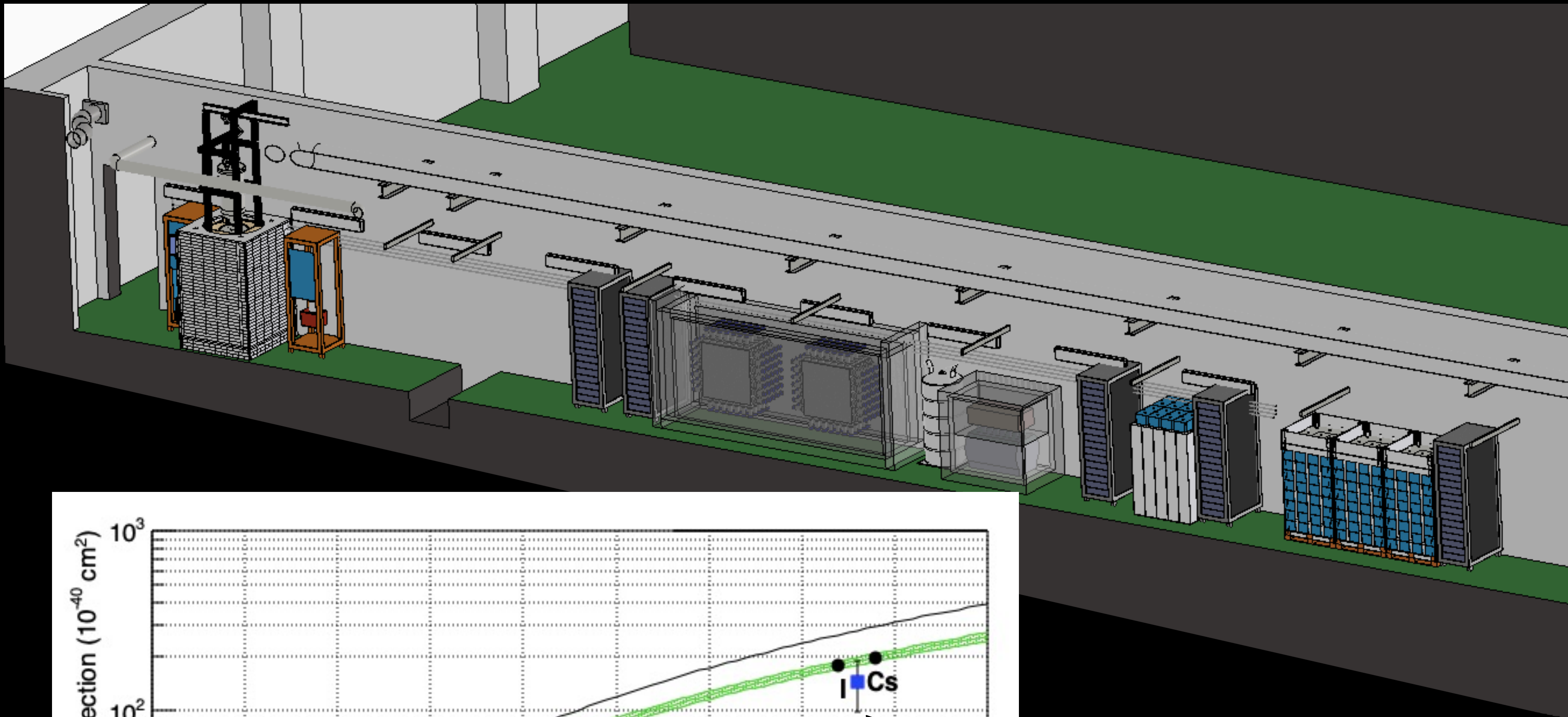
Neutrino non-standard interaction constraints for current CsI data set:

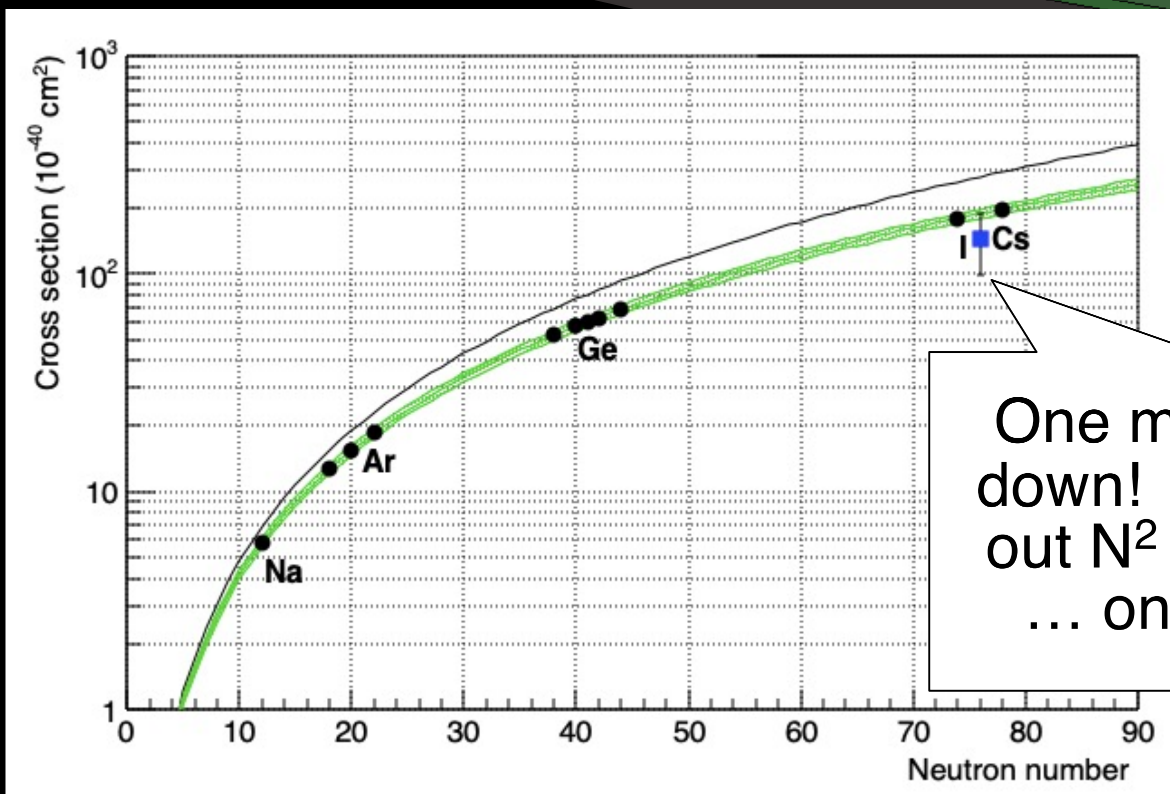


- Assume all other ε's zero

See also Coloma et al., arXiv:1708.02899, many more!

*CHARM constraints apply only to heavy mediators

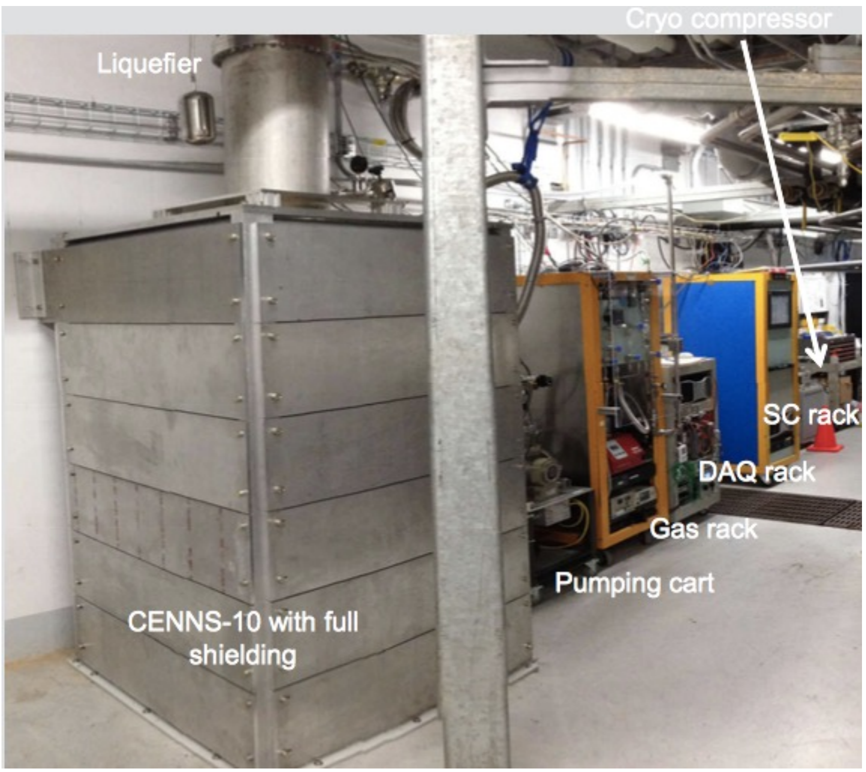




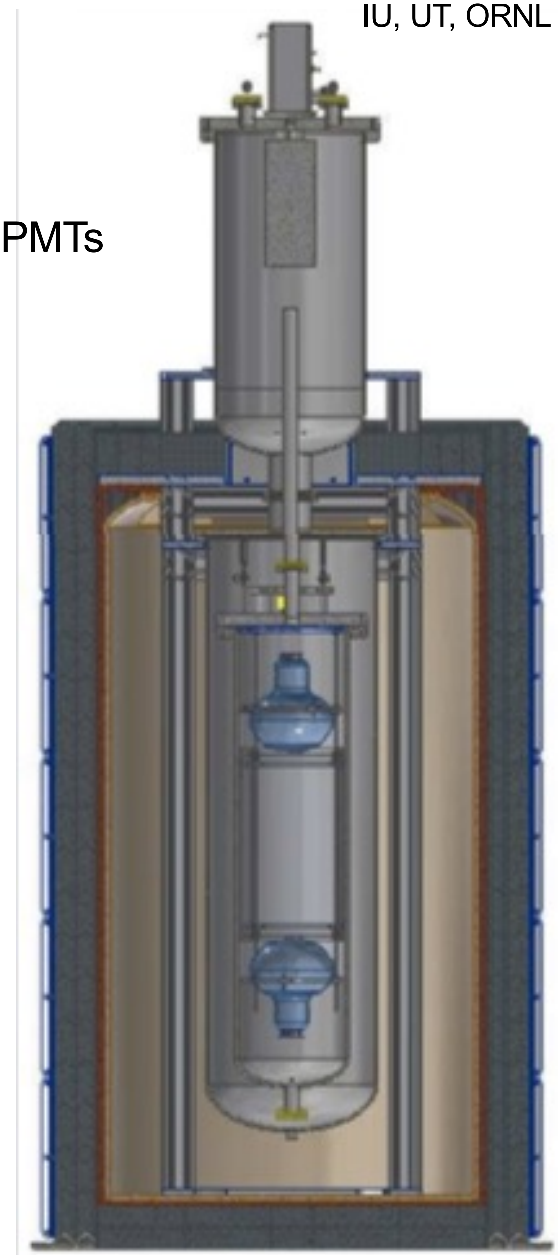
One measurement down! Want to map out N^2 dependence … on to the next

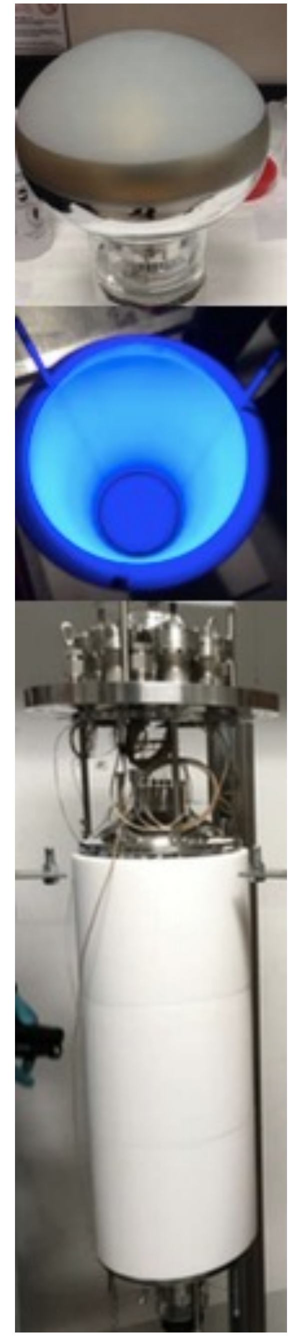
Single-Phase Liquid Argon

- ~24 kg active mass
- 2 x Hamamatsu 5912-02-MOD 8” PMTs
  - 8” borosilicate glass window
  - 14 dynodes
  - QE: 18%@ 400 nm
- Wavelength shifter: TPB-coated Teflon walls and PMTs
- Cryomech cryocooler – 90 Wt
  - PT90 single-state pulse-tube cold head



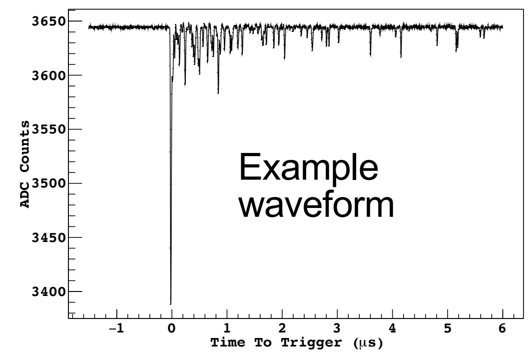
IU, UT, ORNL



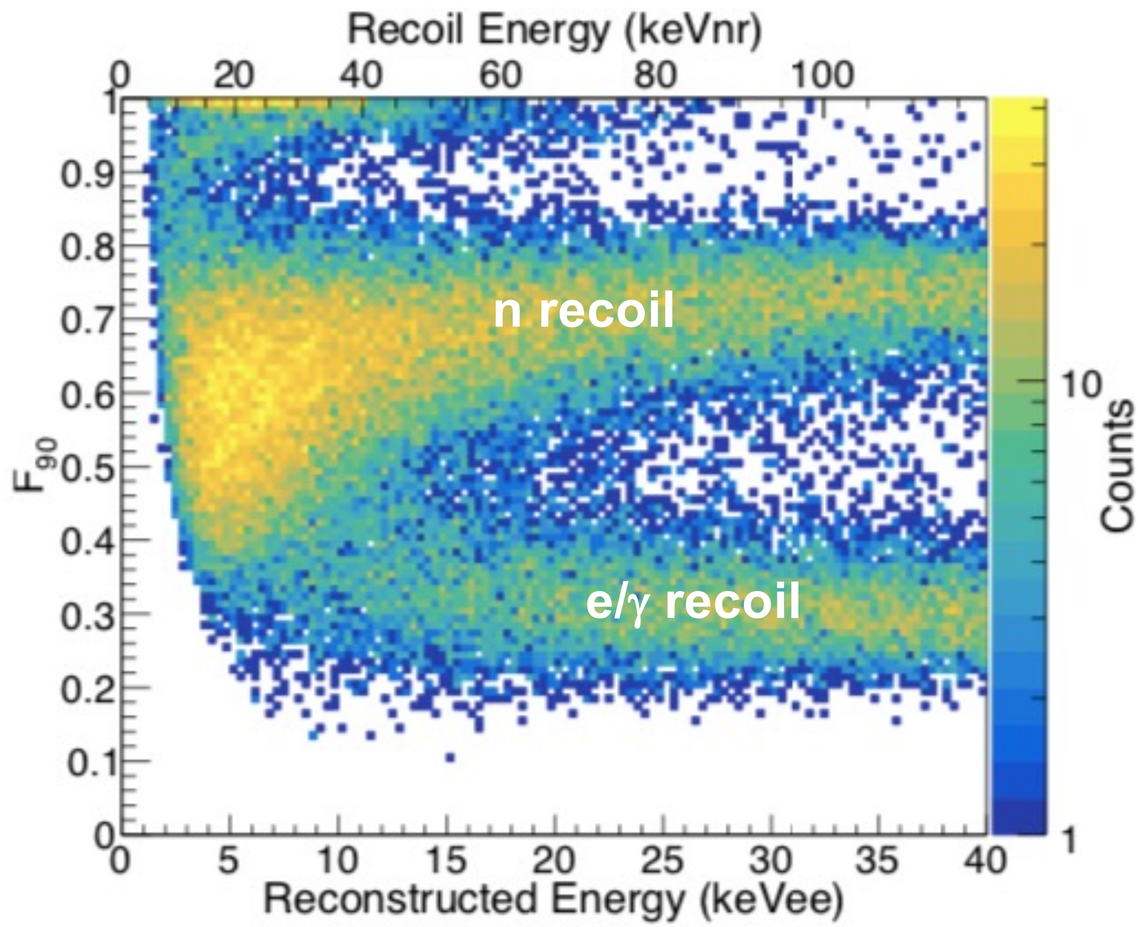


Detector from FNAL, previously built (J. Yoo et al.) for CENNS@BNB
(S. Brice, Phys.Rev. D89 (2014) no.7, 072004)

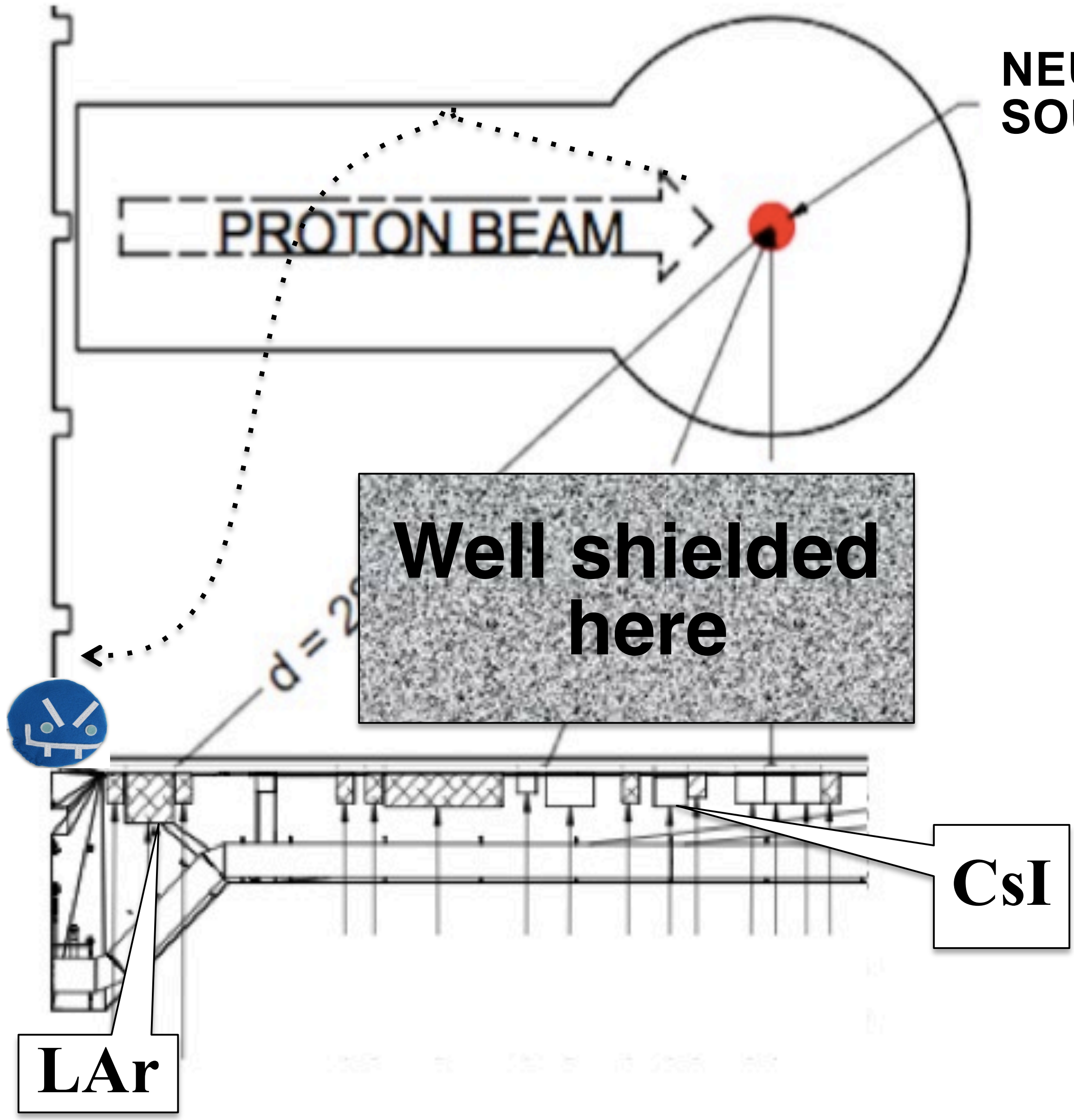
Use **pulse-shape discrimination** to select recoils



F90: fraction of light in first 90 ns



Beam-related neutrons: in the alcove,
need more attention (still tractable)

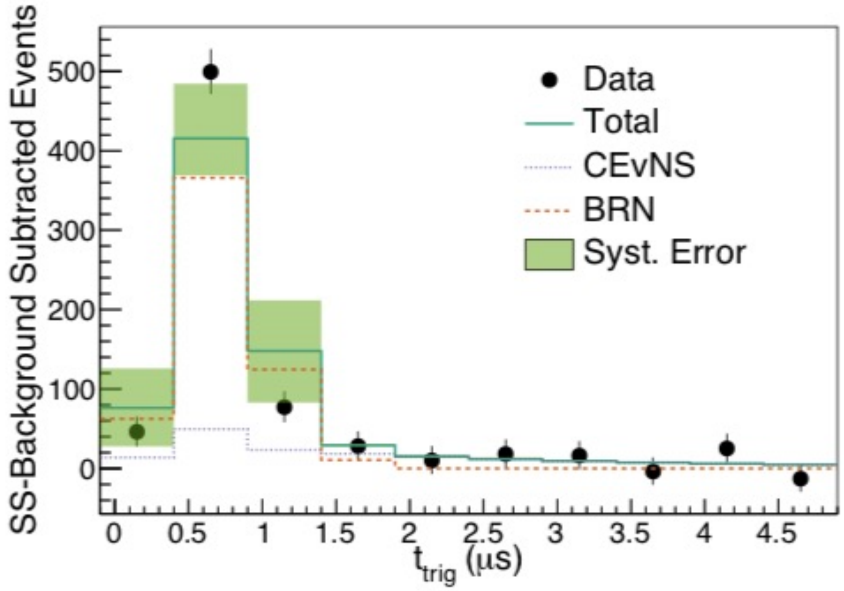


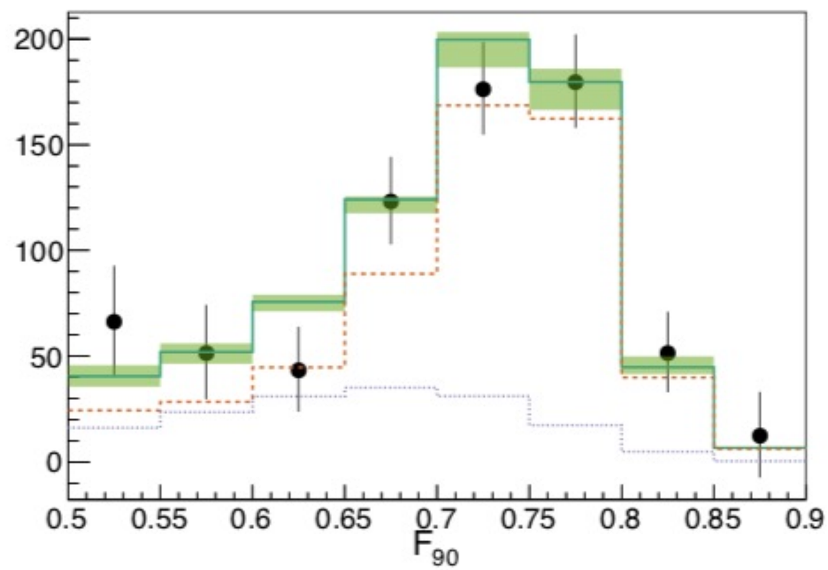
Understand spectrum and time structure by MC tuned using

- Engineering run data
  Phys.Rev. D100 (2019) no.11, 115020
- No-water shield run
- High-energy sideband

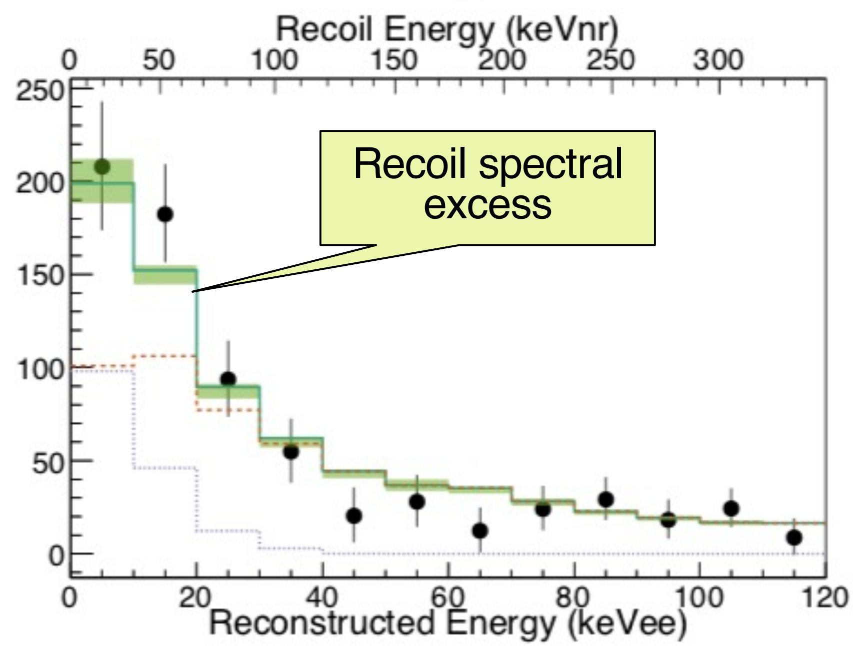
Likelihood fit in time, recoil energy, PSD parameter

Beam-unrelated-background-subtracted projections of 3D likelihood fit

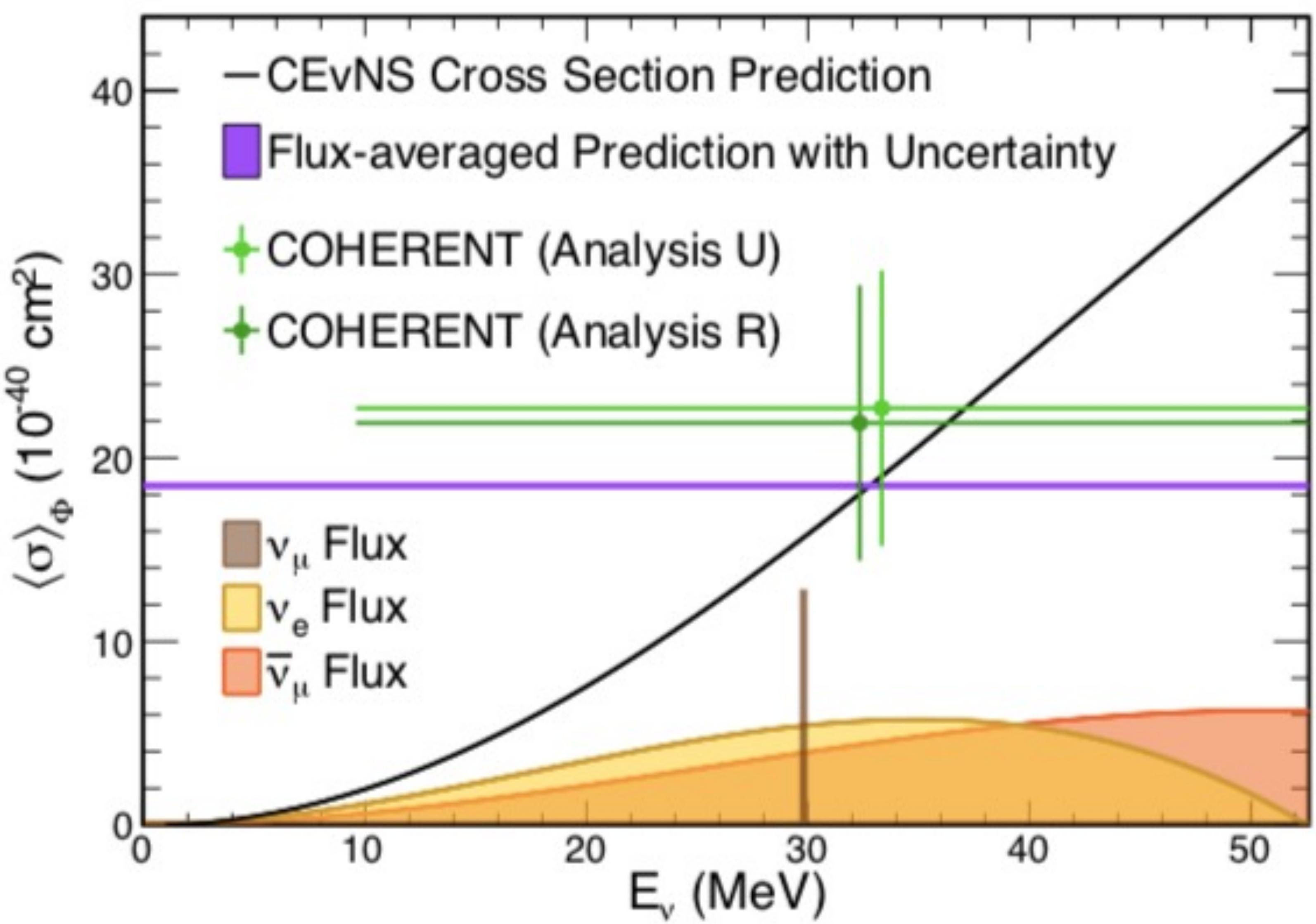




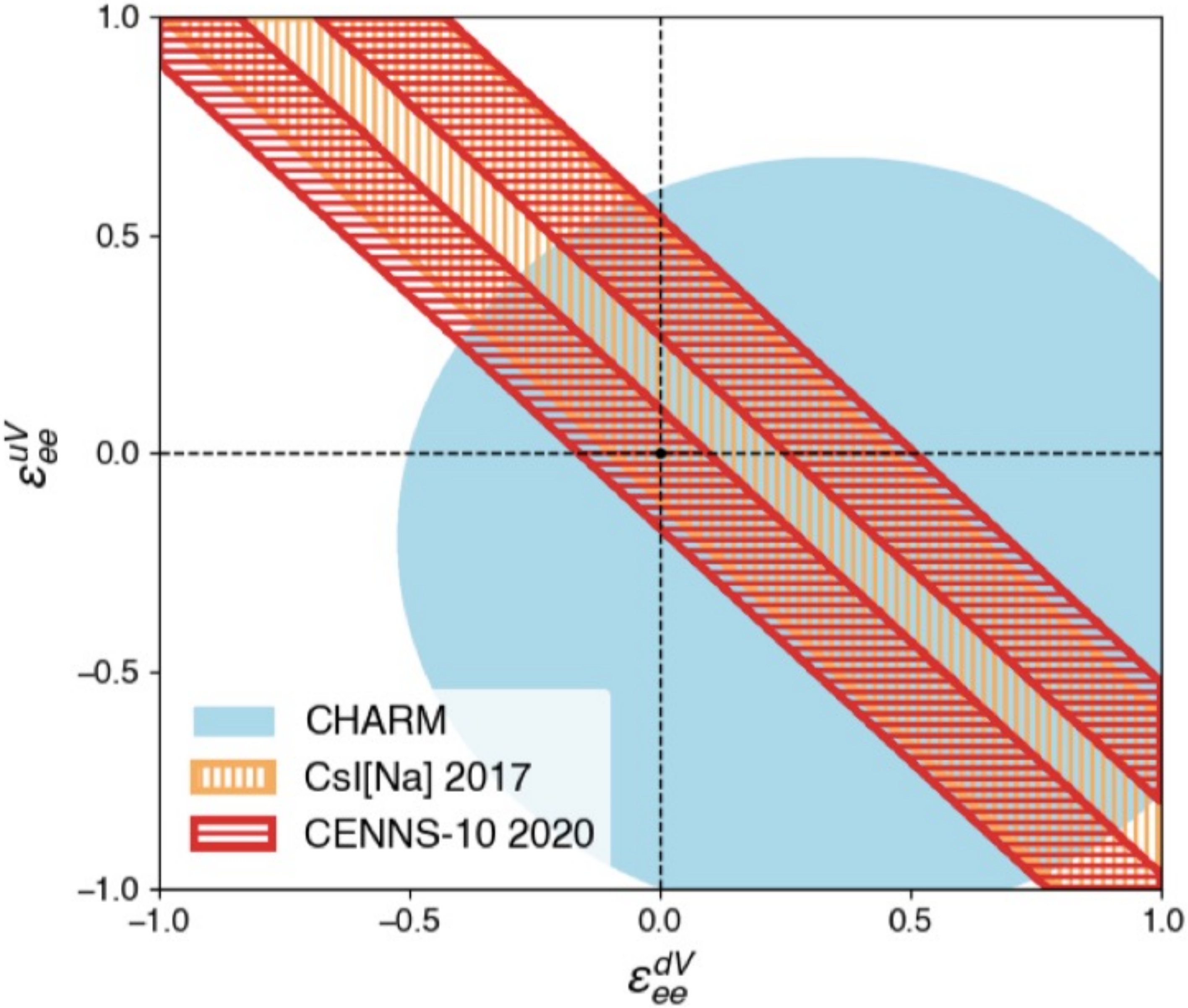
- Bands are systematic errors from 1D excursions
- 2 independent analyses w/separate cuts, similar results (this is the “A” analysis)



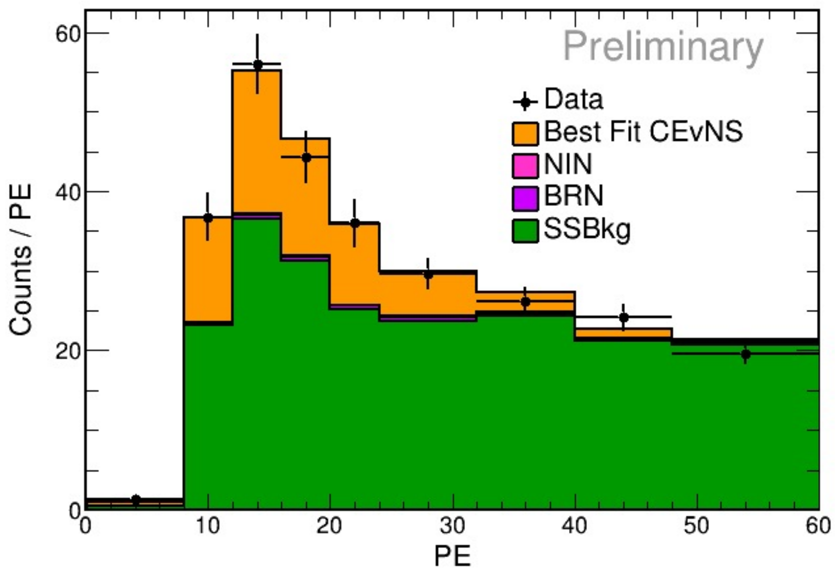
Flux-averaged cross section results

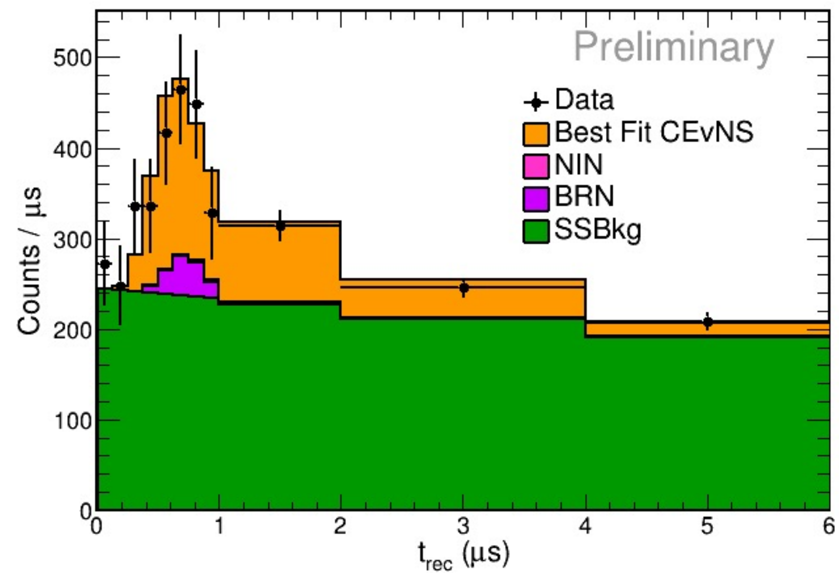


New Constraints on NSI parameters



Remaining CsI[Na] dataset,
with >2 x statistics
+ improved detector response understanding
+ improved analysis





Best fit results

Steady-state background	1273
Beam-related neutrons	17
Neutrino-induced neutrons	5
CEvNS	306

D. Pershey, FNAL JETP seminar

What's the source of the biggest uncertainty on the CEvNS rate for the full COHERENT CsI result?

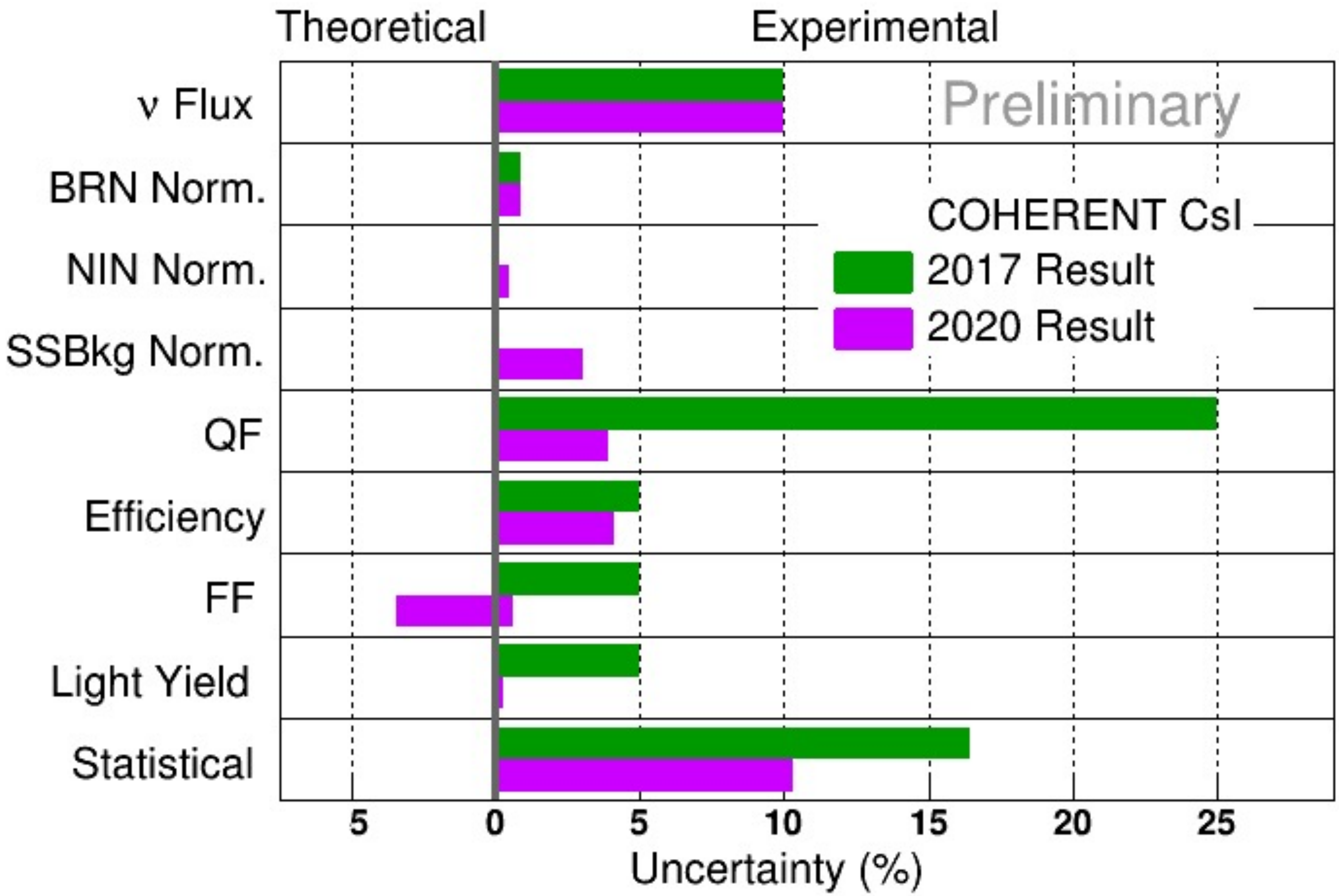
A. Event selection

B. Neutrino flux normalization

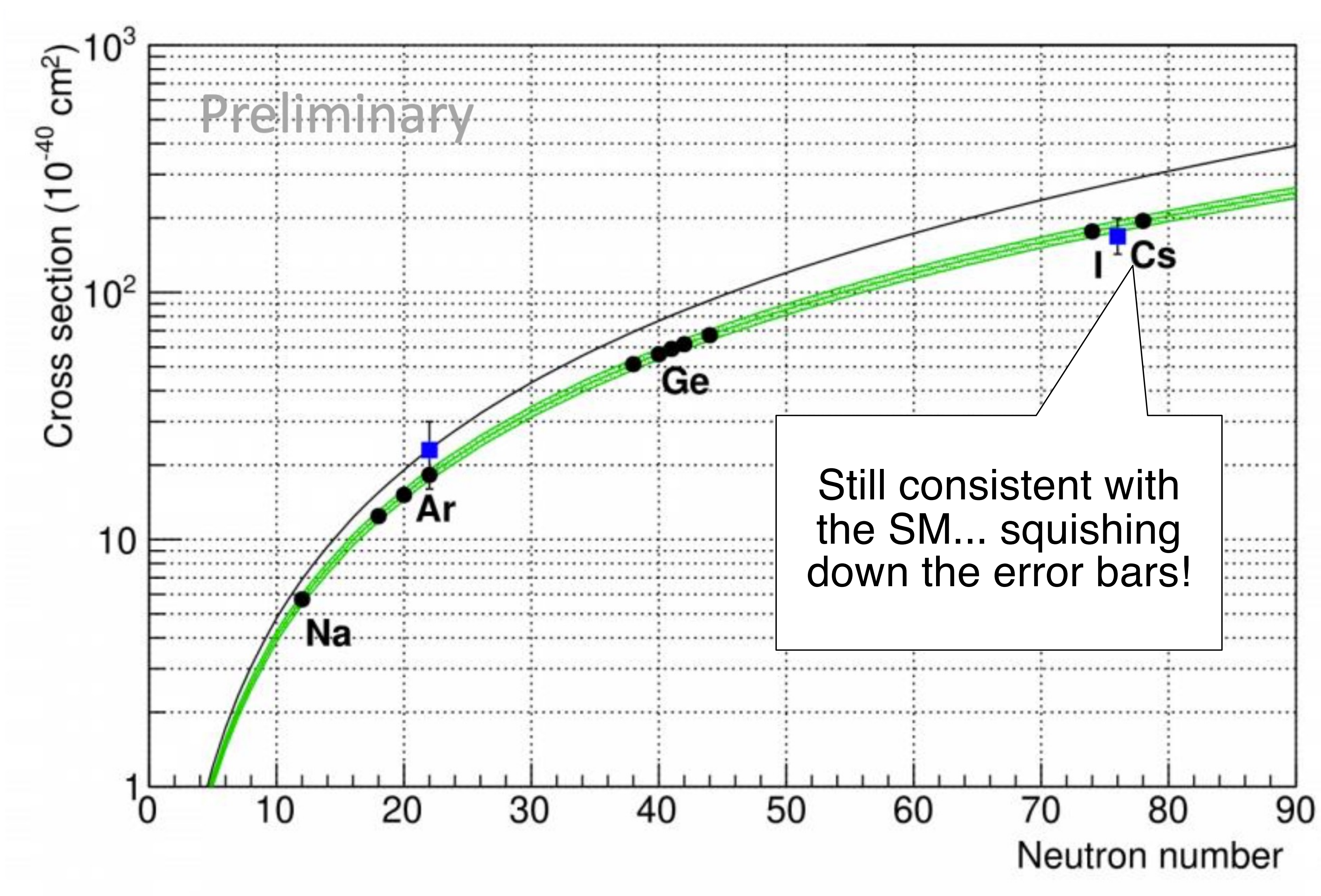
C. Detector response

D. Nuclear structure uncertainties

E. Radiative corrections in the SM

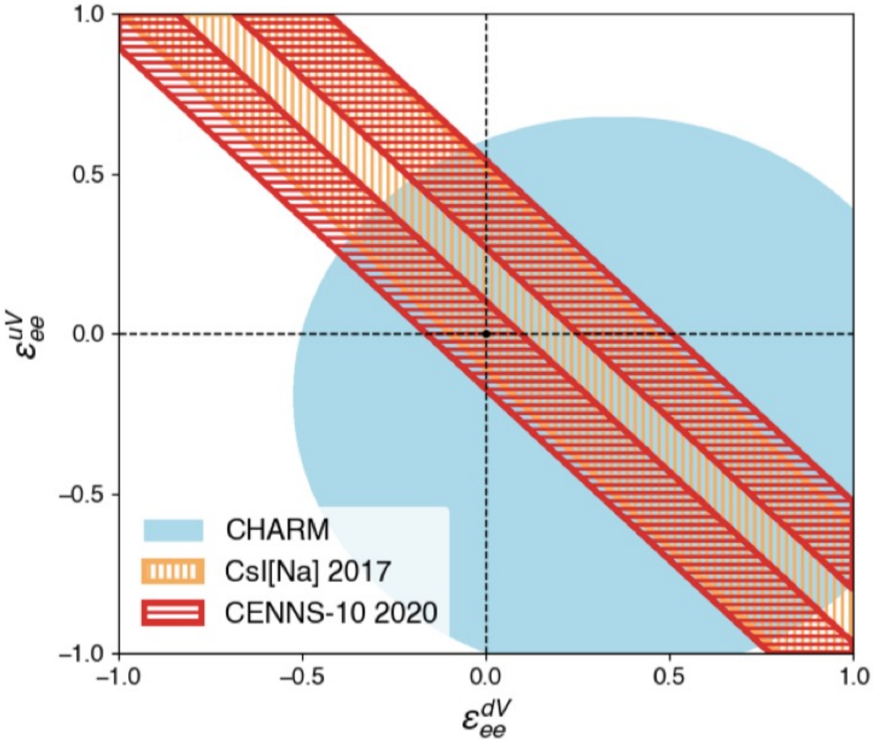


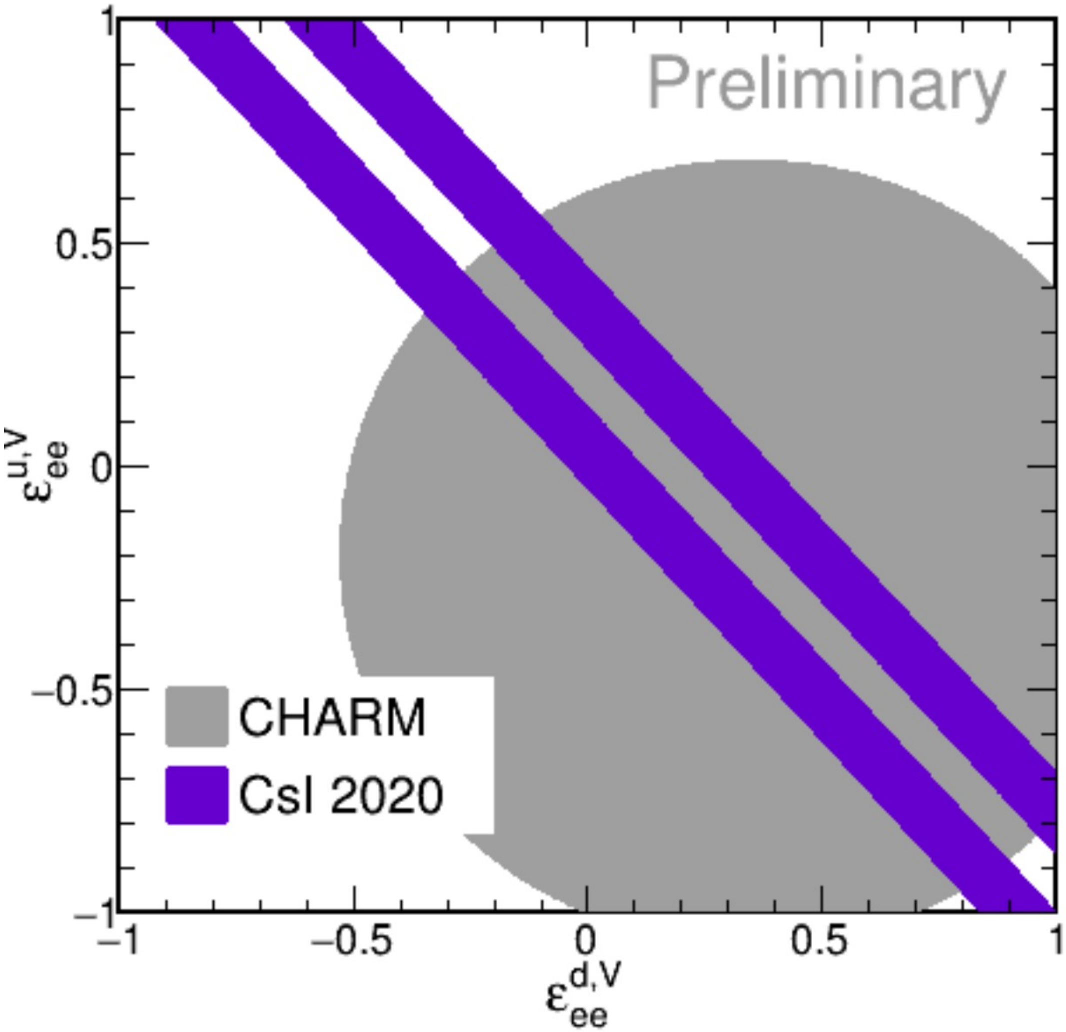
D. Pershey, Magnificent CEvNS talk



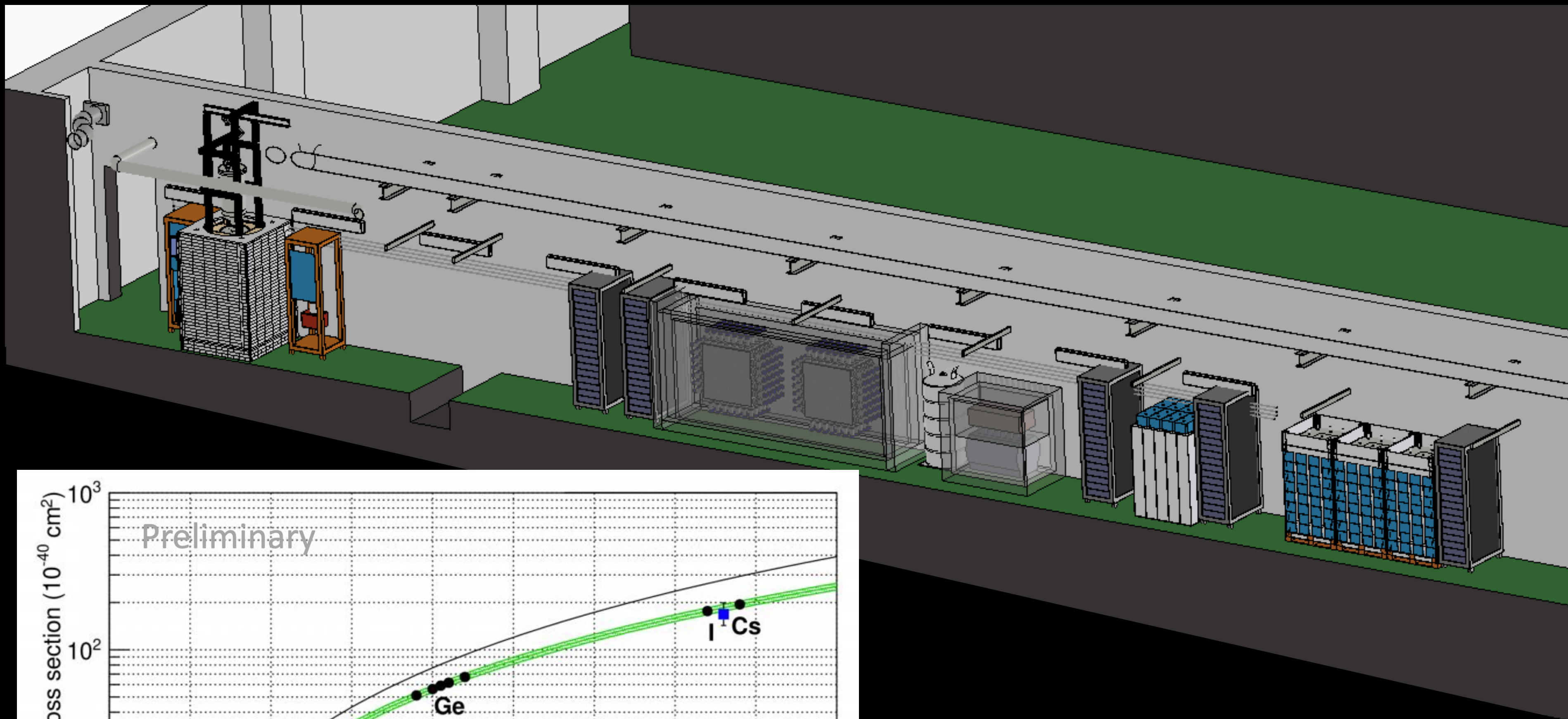
Preliminary
Cross section (10^{-40} cm^2)
10^3
10^2
10
1
0
10
20
30
40
50
60
70
80
90
Neutron number
Na
Ar
Ge
I
Cs
Still consistent with the SM... squishing down the error bars!

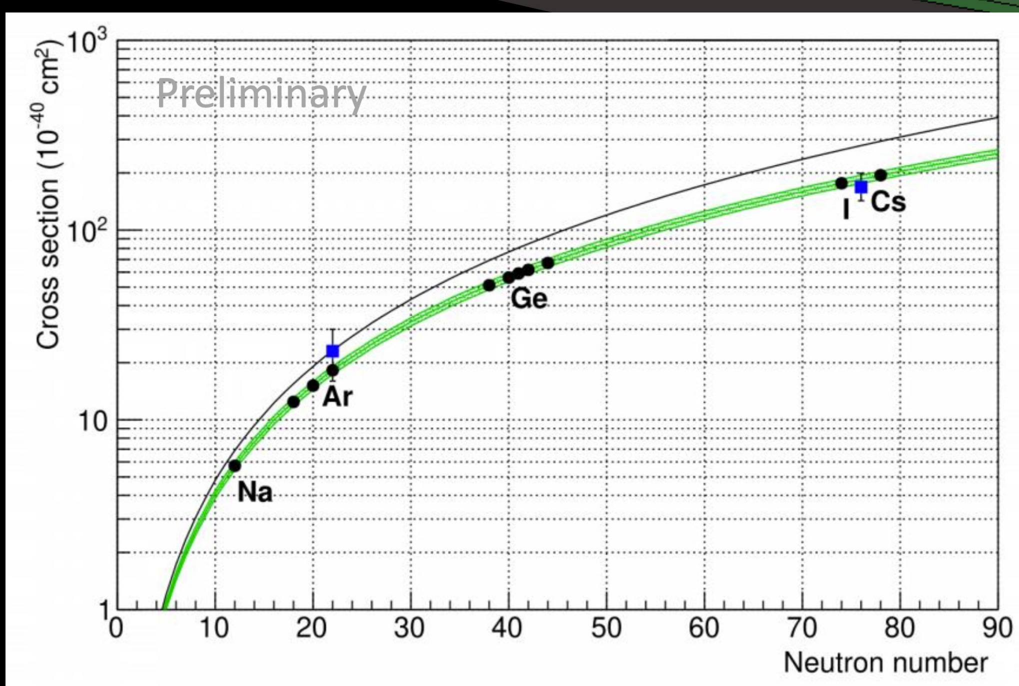
And squeezing down the possibilities for new physics...





What's Next for COHERENT?

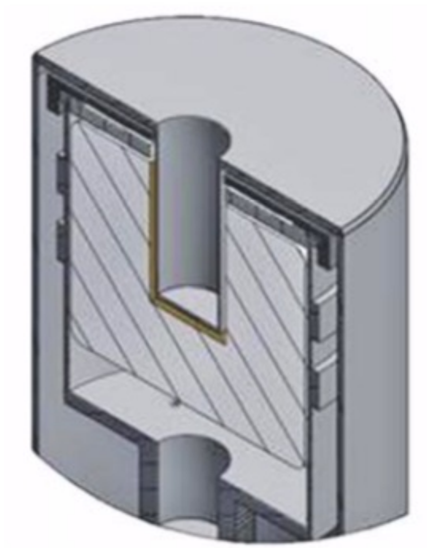




Two down!
But still more to go!

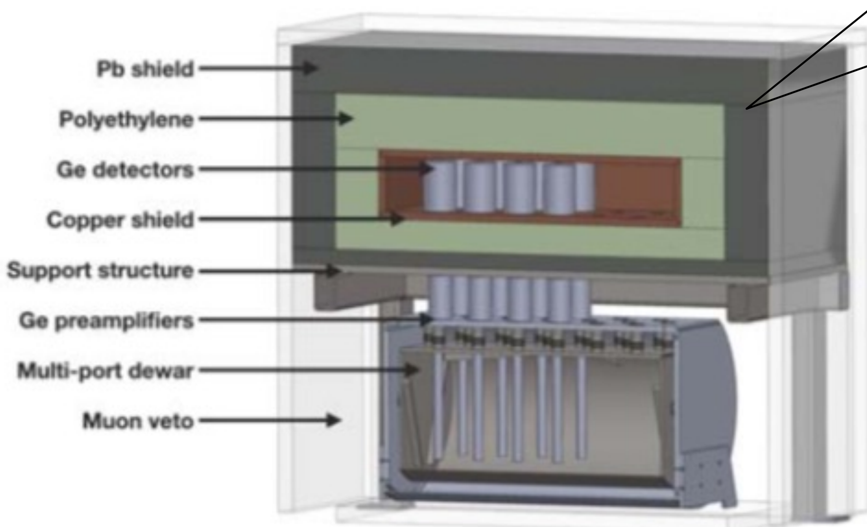
High-Purity Germanium Detectors

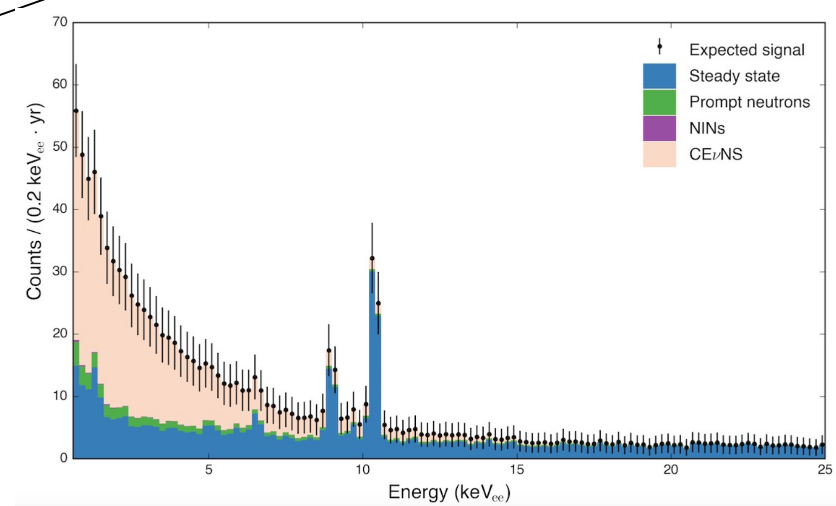
P-type Point Contact



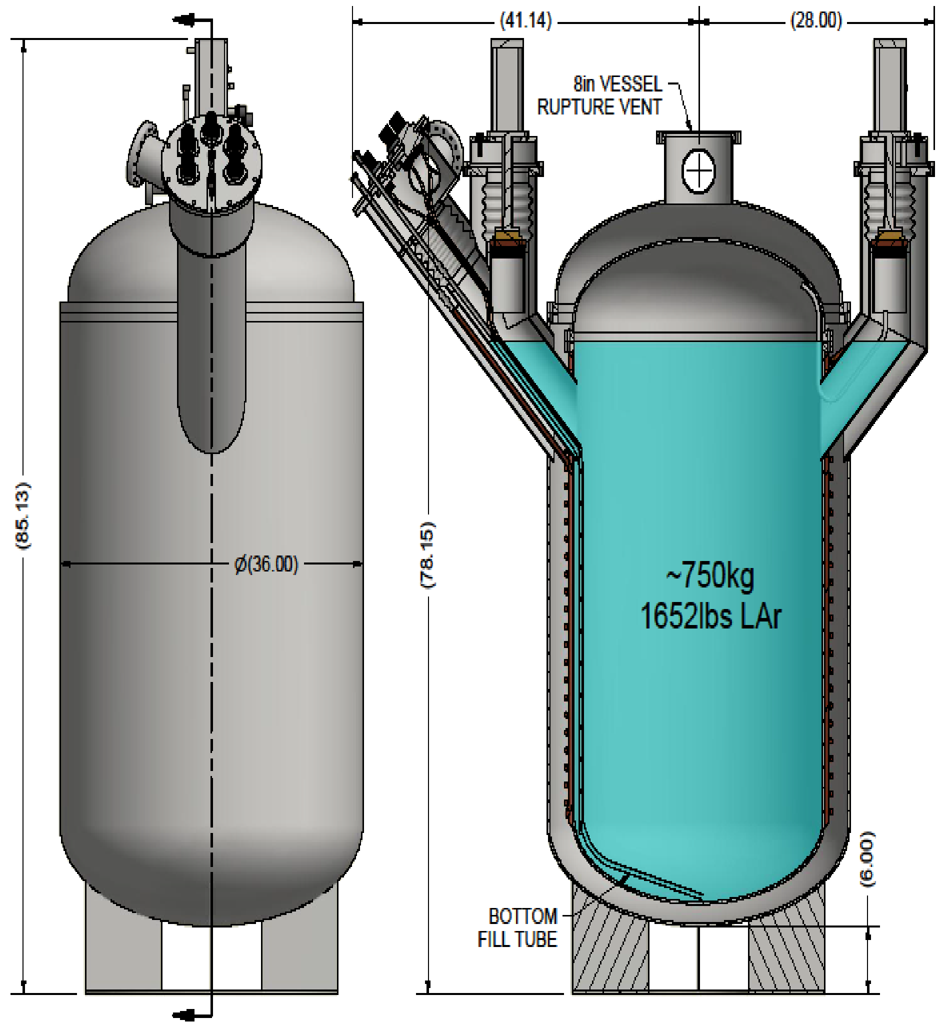
- Excellent low-energy resolution
- Well-measured quenching factor
- Reasonable timing

- 8 Canberra/Mirion 2 kg detectors in multi-port dewar
- Compact poly+Cu+Pb shield
- Muon veto
- Designed to enable additional detectors

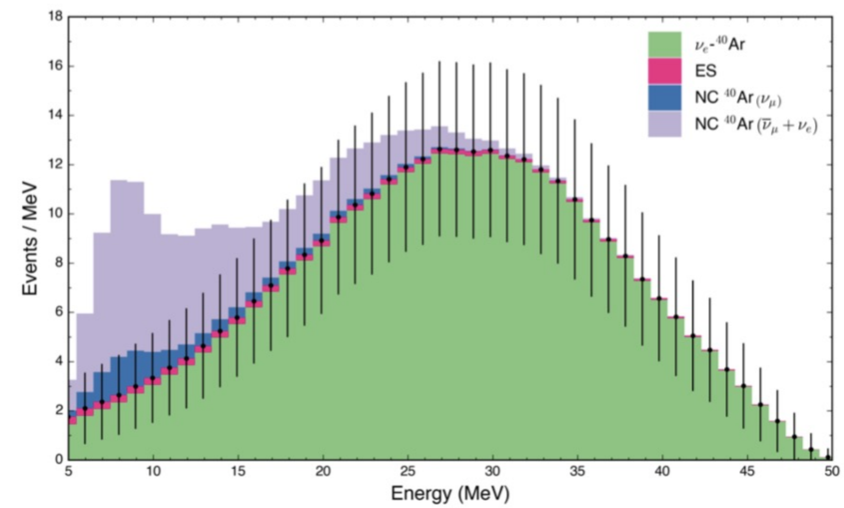




Tonne-scale LAr Detector



- 750-kg LAr will fit in the same place, will reuse part of existing infrastructure
- Could potentially use underground argon



CC/NC **inelastic** in argon of interest for supernova neutrinos

CC $\nu_e + {}^{40}Ar \rightarrow e^- + {}^{40}K^*$

NC $\nu_x + {}^{40}Ar \rightarrow \nu_x + {}^{40}Ar^*$

If ~7000 CEvNS interactions per year are detected in 1 ton of argon at the SNS, about how many ν_eCC events would be expected?

A. 70000

B. 3000

C. 300

D. 70

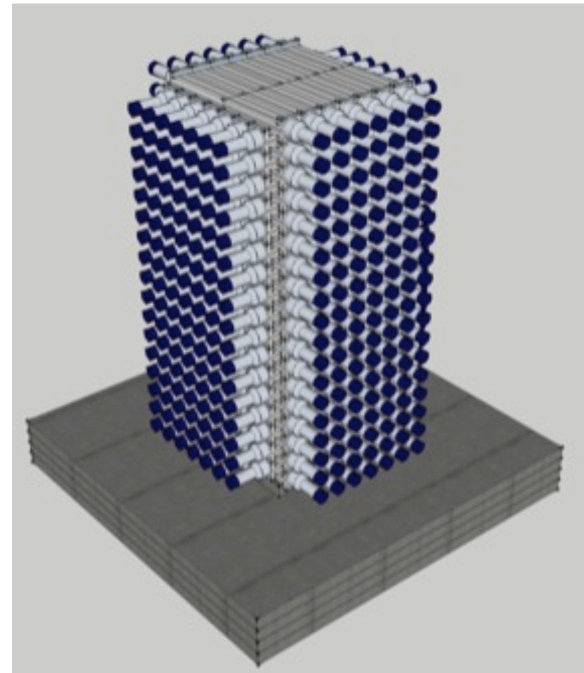
E. 3

F. 0.1

Sodium Iodide (NaI[Tl]) Detectors (NaIvE)

- up to 9 tons available, 2 tons in hand
- QF measured
- require PMT base refurbishment (dual gain) to enable low threshold for CEvNS on Na measurement
- development and instrumentation tests underway at UW, Duke

Multi-ton concept



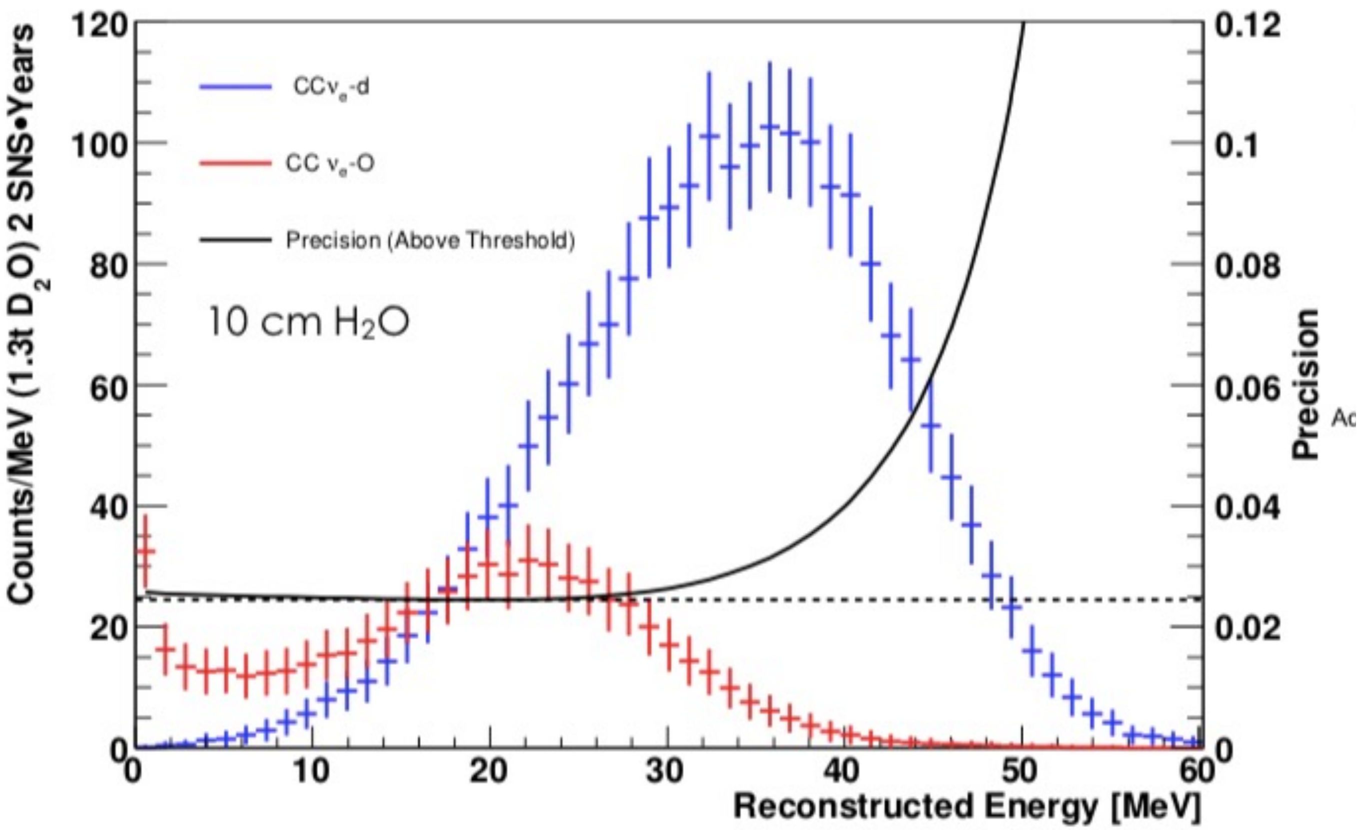
In the meantime: **185 kg deployed at SNS** to go after ν_e**CC on ^{127}I**

Isotope	Reaction Channel	Source	Experiment	Measurement (10^{-42} cm^2)	Theory (10^{-42} cm^2)
^{127}I	^{127}I$(\nu_e, e^-)^{127}$Xe	Stopped π/μ	LSND	284 ± 91(stat) ± 25(sys)	210-310 [Quasi-particle] (Engel *et al.*, 1994)

J.A. Formaggio and G. Zeller, RMP 84 (2012) 1307-1341

Heavy water detector in Neutrino Alley

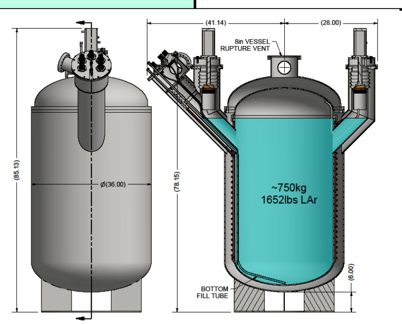
Measurement Precision with 2 SNS years at 1.4 MW

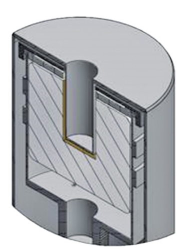


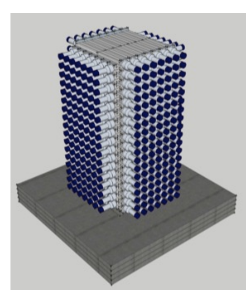
➔ ~few percent precision on flux normalization

COHERENT **CEvNS** Detector Status and **Farther Future**

Nuclear Target	**Technology**	**Mass (kg)**	**Distance from source (m)**	**Recoil threshold (keVr)**	**Data-taking start date**	**Future**
CsI[Na]	Scintillating crystal	14.6	19.3	6.5	9/2015	Decommissioned
Ge	HPGe PPC	18	22	<few	2021	Funded by NSF MRI, in progress
LAr	Single-phase	24	27.5	20	12/2016, upgraded summer 2017	Expansion to **750 kg scale**
NaI[Tl]	Scintillating crystal	185*/ 3388	25	13	*high-threshold deployment summer 2016	Expansion to **3.3 tonne**, up to 9 tonnes



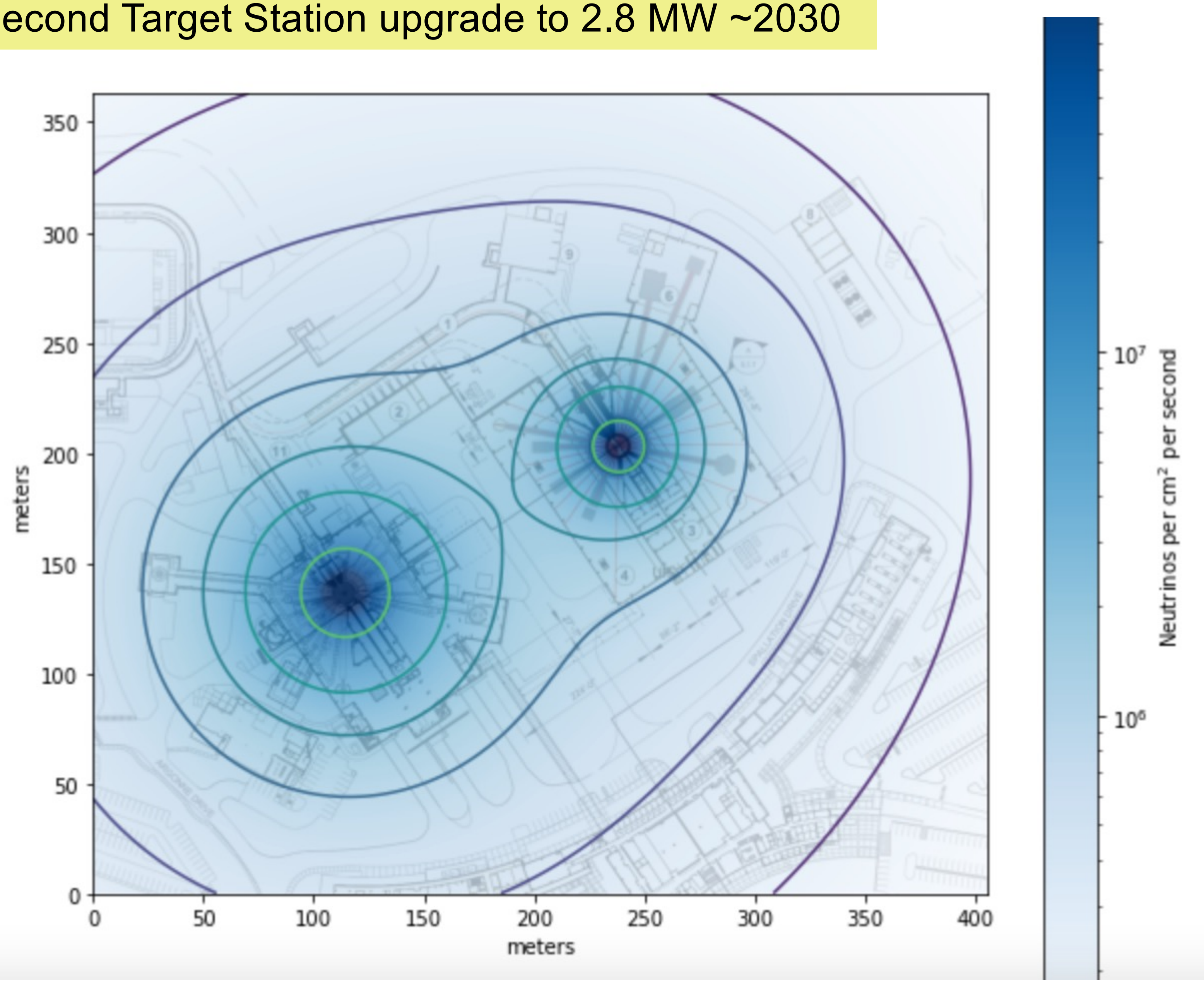




+D_2O for flux normalization
+ concepts for other targets...

+ power upgrade to 2 MW in 2023, Second Target Station upgrade to 2.8 MW ~2030

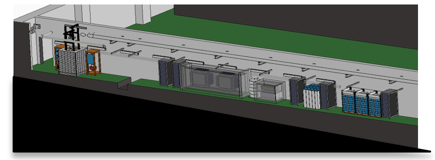
SNS power upgrade to 2 MW in 2023,
Second Target Station upgrade to 2.8 MW ~2030

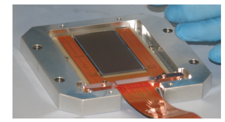
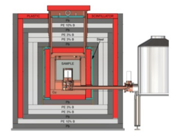


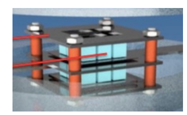
Many exciting possibilities for ν's + DM!

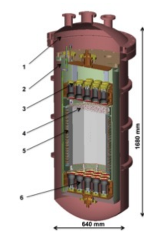
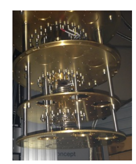
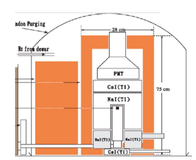
Many CEvNS Efforts Worldwide

Experiment	Technology	Location	Source
COHERENT	CsI, Ar, Ge, NaI	USA	πDAR
CCM	Ar	USA	πDAR
CONNIE	Si CCDs	Brazil	Reactor
CONUS	HPGe	Germany	Reactor
MINER	Ge/Si cryogenic	USA	Reactor
NuCleus	Cryogenic $CaWO_4$, Al_2O_3 calorimeter array	Europe	Reactor
νGEN	Ge PPC	Russia	Reactor
RED-100	LXe dual phase	Russia	Reactor
Ricochet	Ge, Zn bolometers	France	Reactor
TEXONO	p-PCGe	Taiwan	Reactors



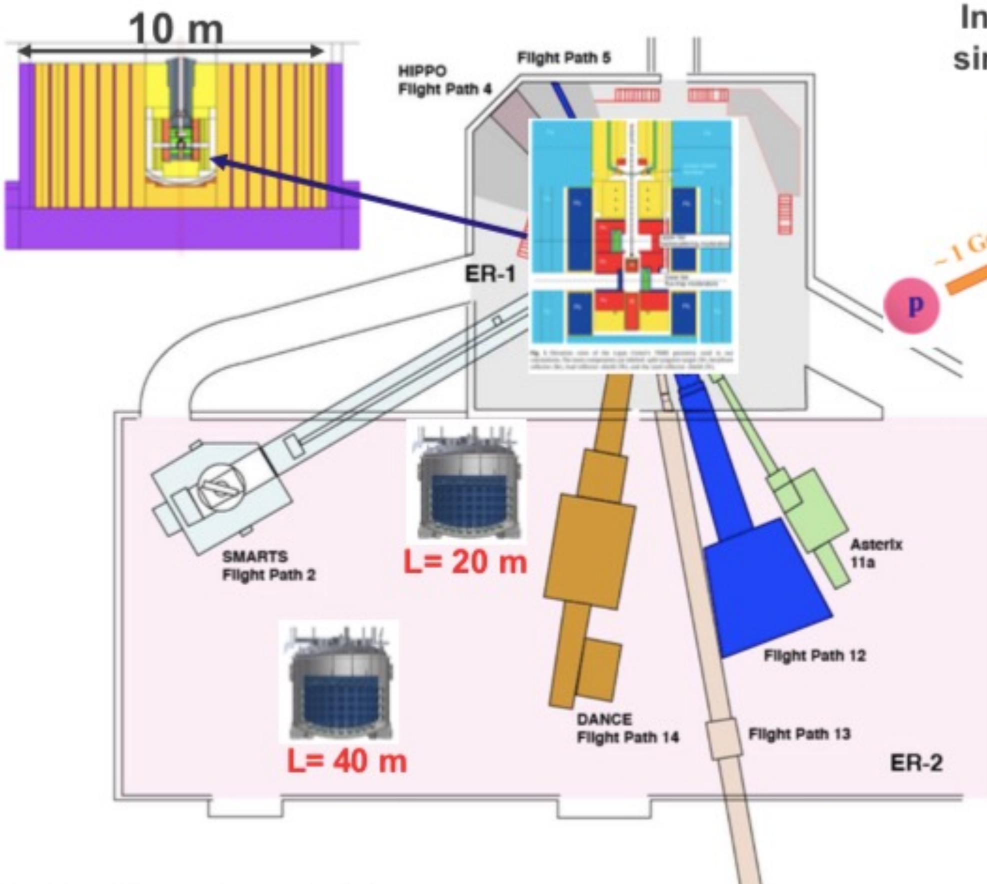


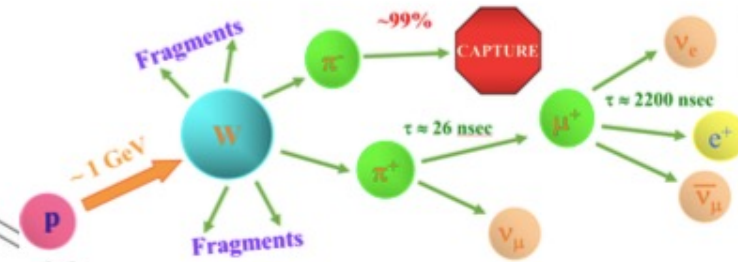
+ DM detectors, +directional detectors +more...

many novel low-background, low-threshold technologies!!

Coherent Captain Mills @ Lujan: single-phase LAr



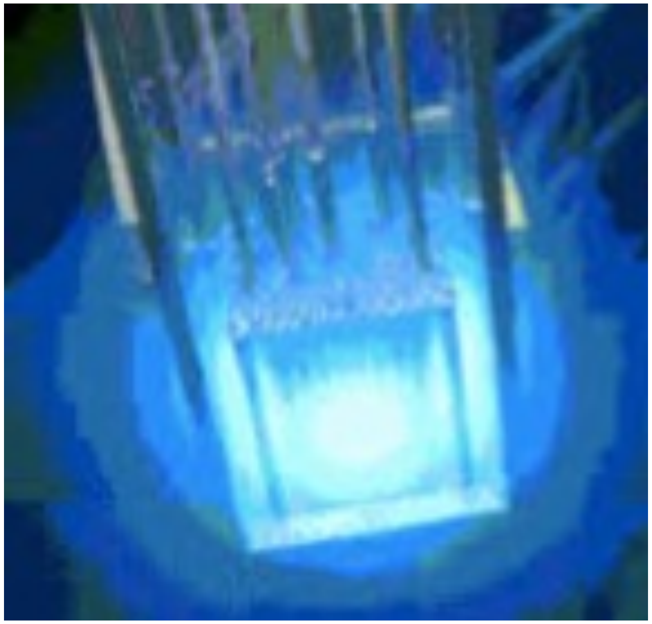
Intense source muon neutrinos: target MCNP simulation flux 4.74x10^5 ν/cm^2/s at 20 m

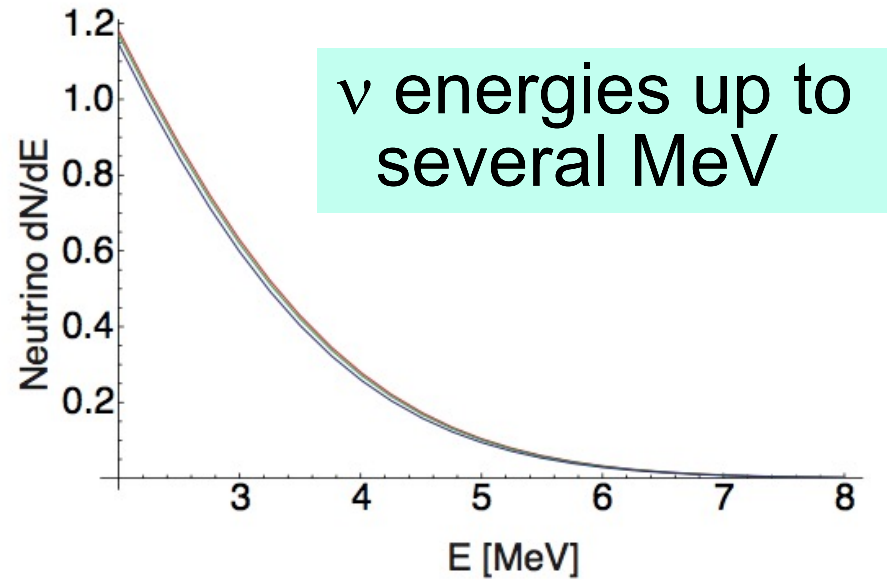


Lujan Experimental Area
- Space for large 10-ton liquid Argon ν detector.
- Run detector in multiple locations.
- Room to deploy shielding, large overhead crane, power, etc

Primary focus on sterile neutrinos & DM search

Neutrinos from nuclear reactors



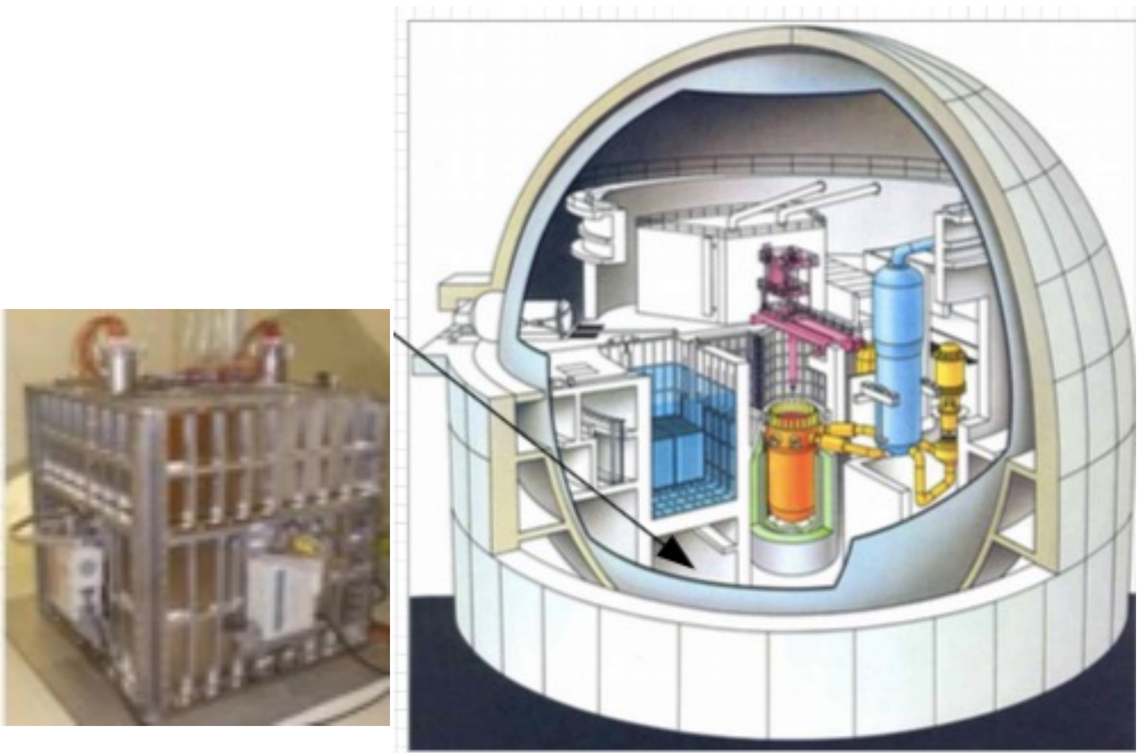


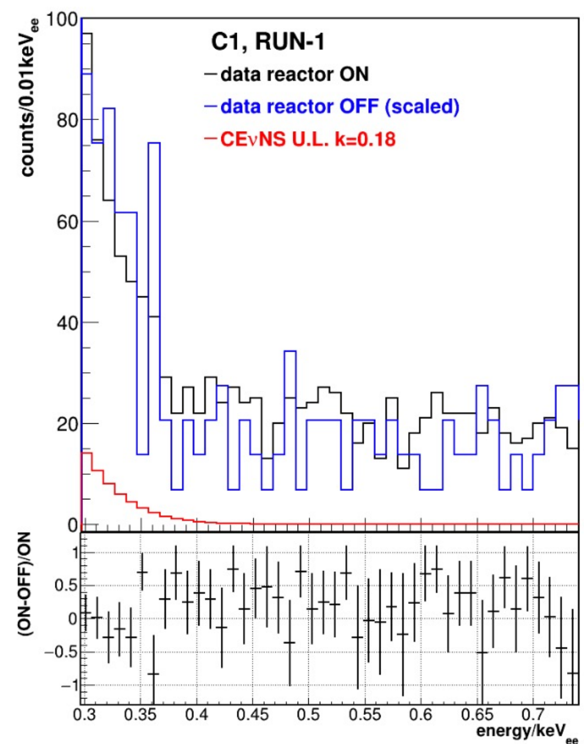
- ν_e-bar produced in fission reactions (one flavor)
- **huge fluxes possible**: ~$2\text{x}10^{20}$ s^{-1} per GW
- several CEvNS searches past, current and future at reactors, but **recoil energies<keV** and backgrounds make this very challenging

CONUS

- Brokdorf 3.9 GWth reactor, Germany
- 17 m from core
- 4 kg Ge PPC
- ~300 eVee threshold

Phys.Rev.Lett. 126 (2021) 4, 041804 arXiv: 2011.00210 [hep-ex]

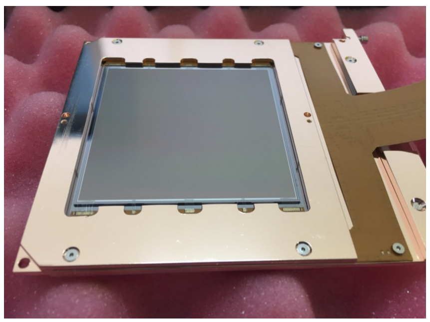


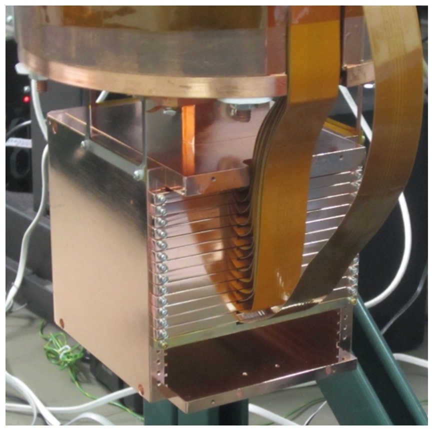


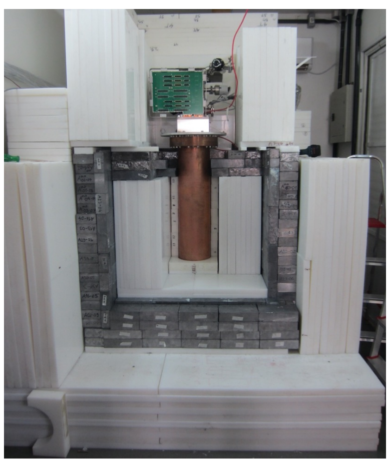
<85 events in ROI @90 CL

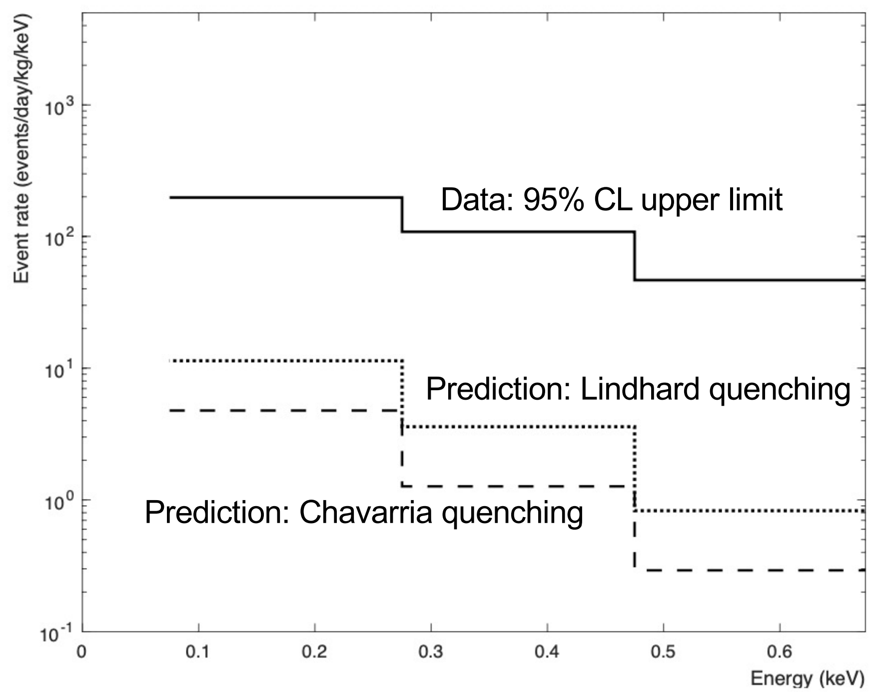
CONNIE

- Angra-2 3.8 GWth nuclear reactor, Brazil
- 32 m from core
- 47.6 g Si CCDs
- ~0.1 keVee threshold



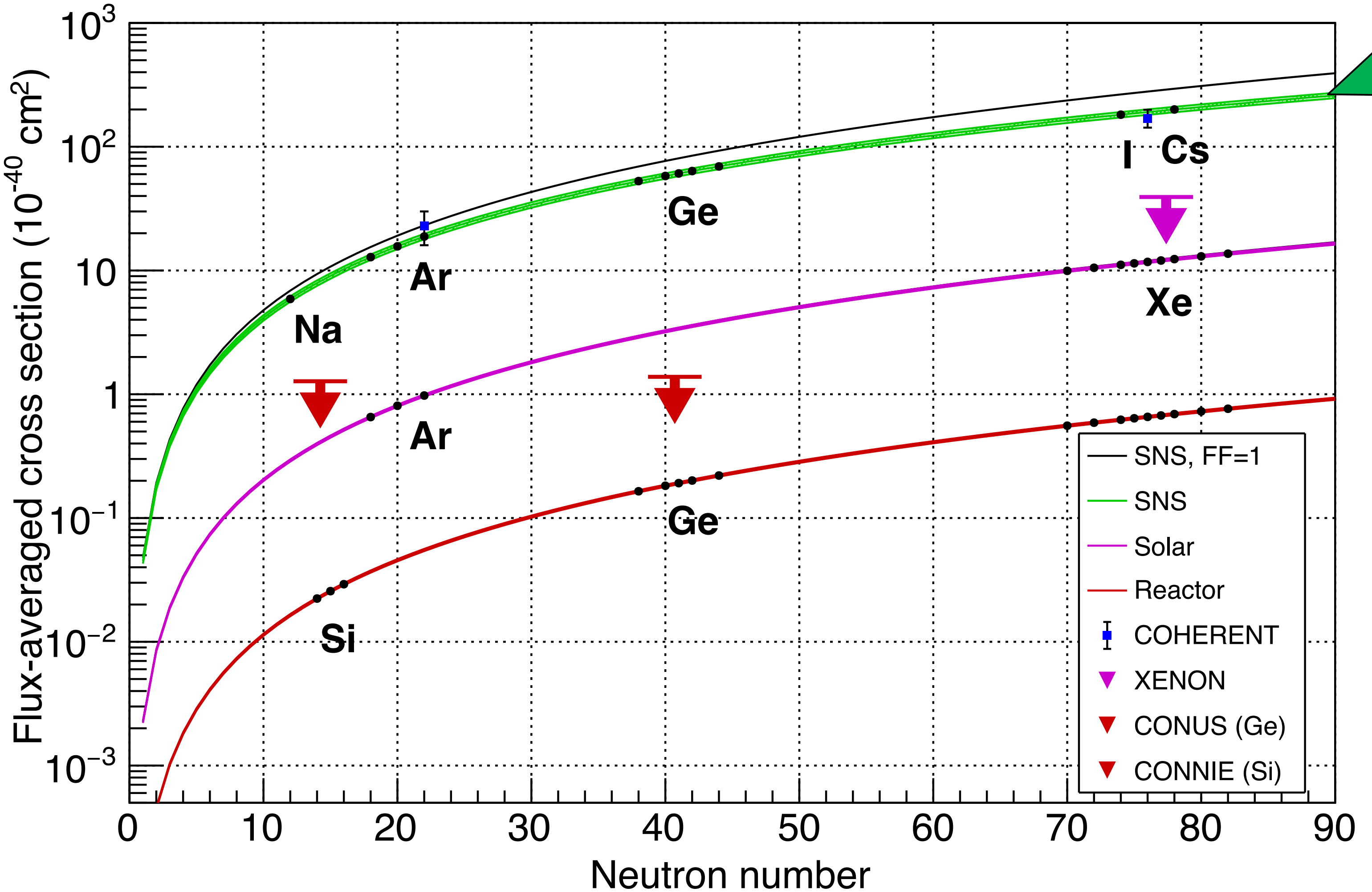






Phys.Rev.D 100 (2019) 9, 092005
arXiv: 1906.02200 [physics.ins-det]

Summary of CEvNS Results



So far: measurements in CsI, Ar from COHERENT
... looking forward to more soon!

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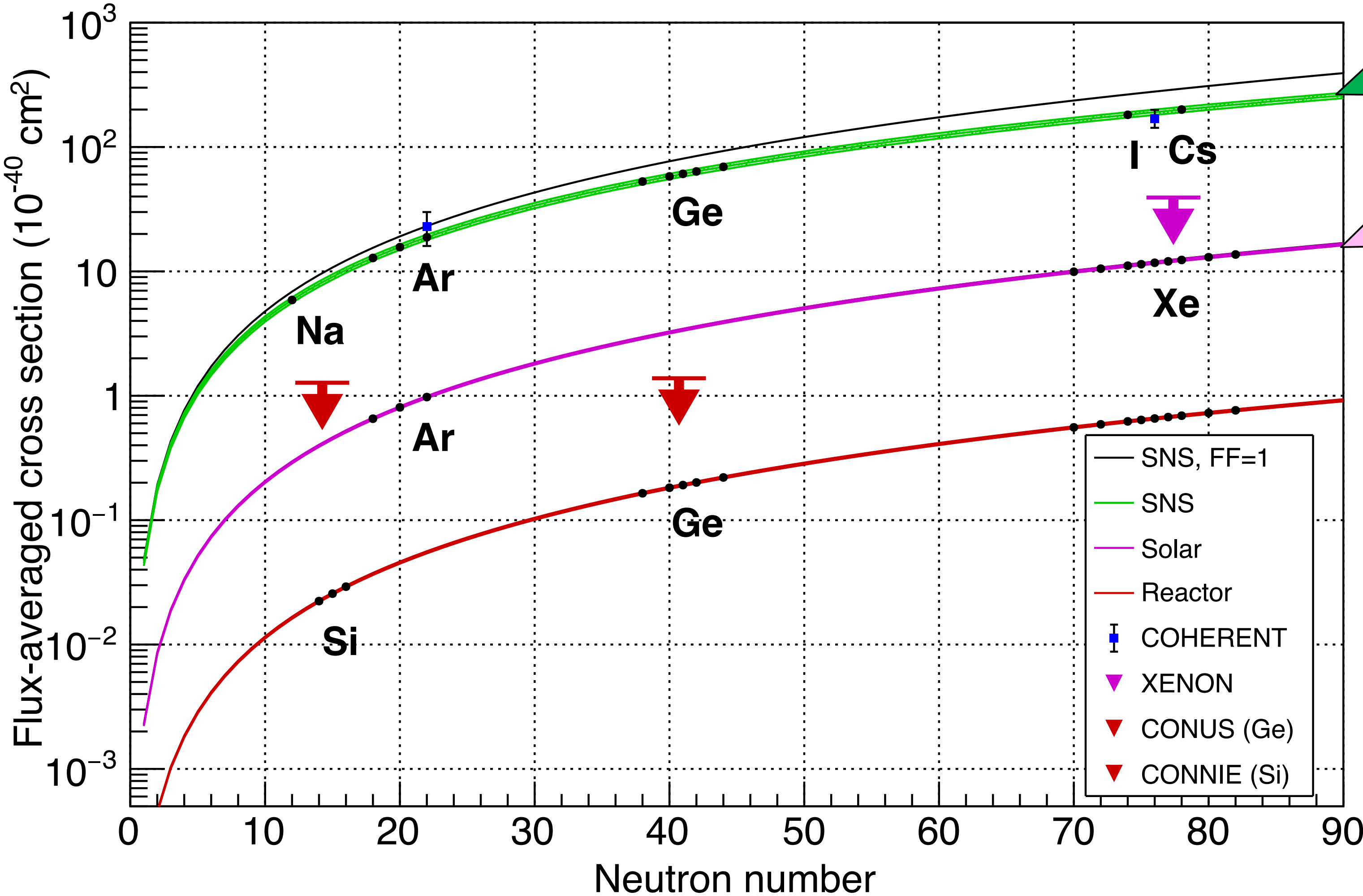

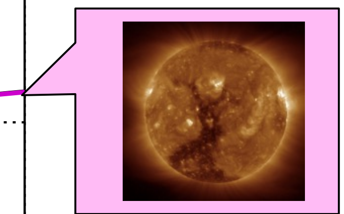

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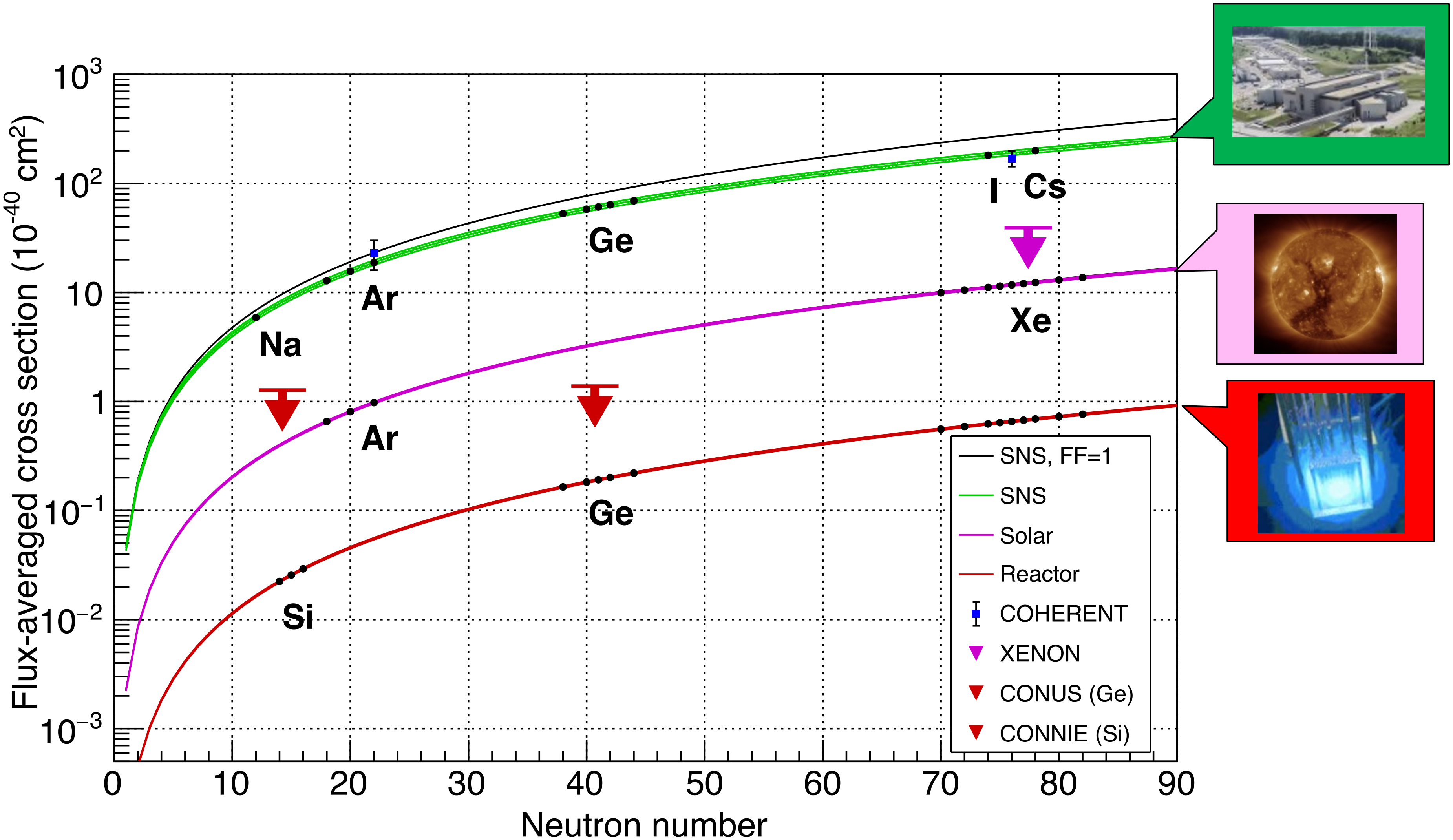
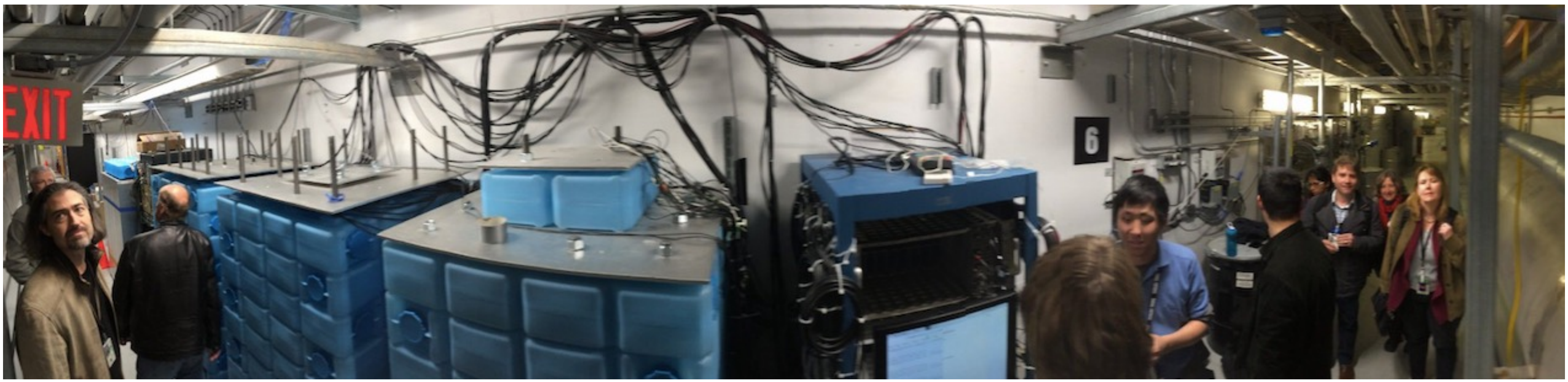
Summary of CEvNS Results



So far: measurements in CsI, Ar from COHERENT
... looking forward to more soon!

Take-away points from the lecture

- **CEvNS:**
  - large cross section, but tiny recoils, α N^2
  - accessible w/low-energy threshold detectors, plus extra oomph of stopped-pion neutrino source
- **First measurement** by COHERENT CsI[Na] at the SNS... now Ar, + more CsI data!
- **Meaningful bounds on beyond-the-SM physics**



- Nore NaI+Ge CEvNS soon, (+ inelastics)!
- Multiple targets, upgrades and new ideas in the works!
- "Neutrino Avenue" at the Second Target Station?

- Other CEvNS experiments will join the fun!
  (CCM, TEXONO, CONUS, CONNIE, MINER, RED, Ricochet, NUCLEUS...)