

Neutrino Self-interaction as a solution(?) to the Hubble Tension

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Based on 2011.12315
with Subhajit Ghosh



Outline

- **Cosmology of self-interacting neutrino**
- **Flavor-universal $S_{\nu\nu}$ & the Hubble tension**
- **Laboratory constraints on $S_{\nu\nu}$**
- **Flavor-specific $S_{\nu\nu}$**

A phenomenological model

$$\mathcal{L} \supset \frac{1}{2} g_{ij} \bar{\nu}_i \nu_j \phi, \quad g_{ij} = g \delta_{ij}$$



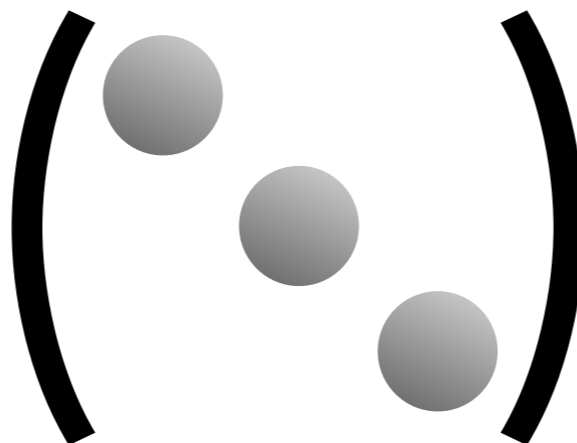
$$\mathcal{L} \supset G_{\text{eff}} \bar{\nu} \nu \bar{\nu} \nu, \quad G_{\text{eff}} = \frac{g^2}{M_\phi^2}$$

1306.1536

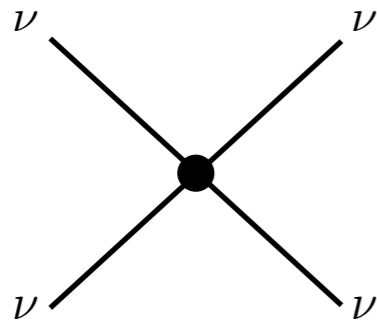
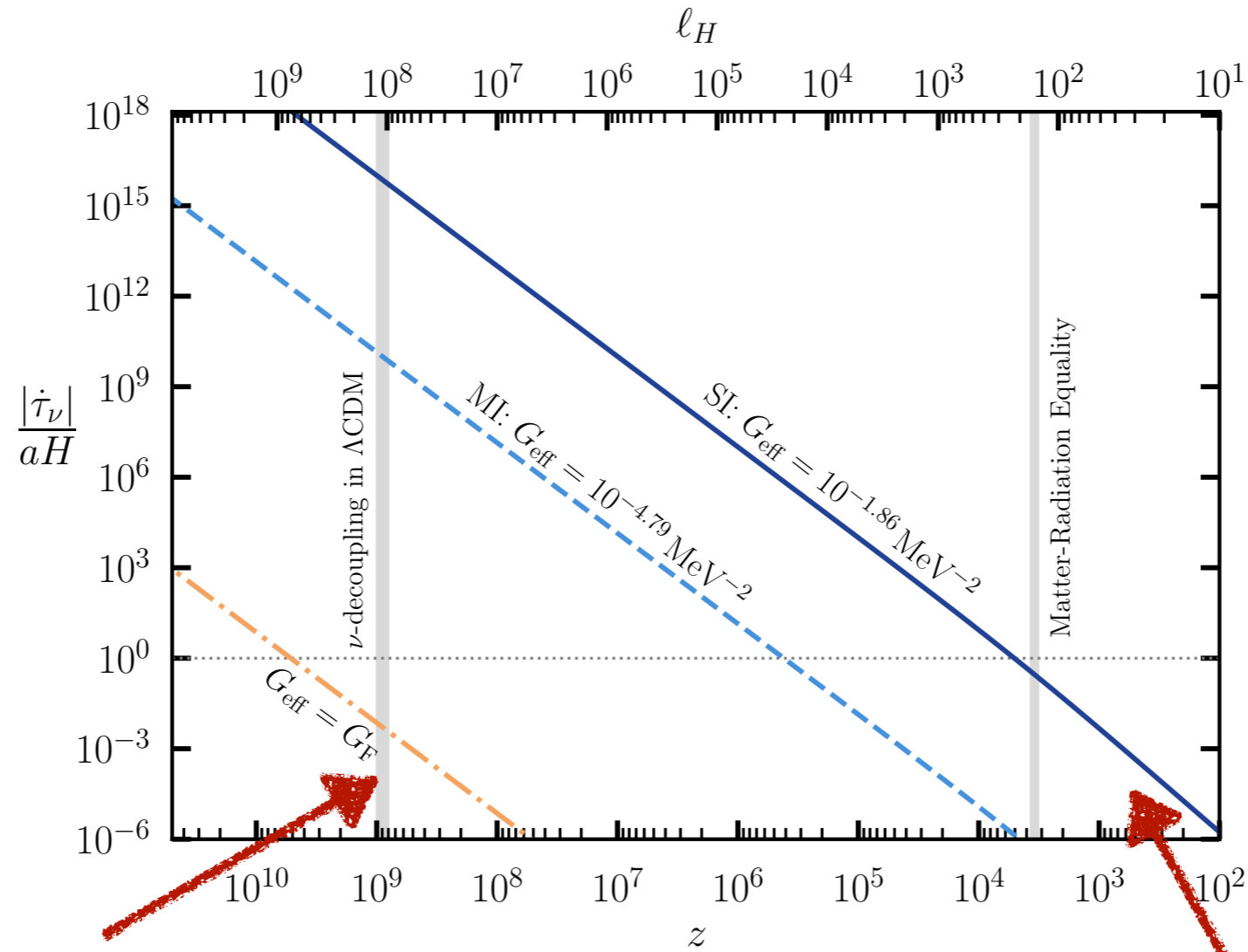
1704.06657

1902.00534

All neutrino flavors interact with the same strength

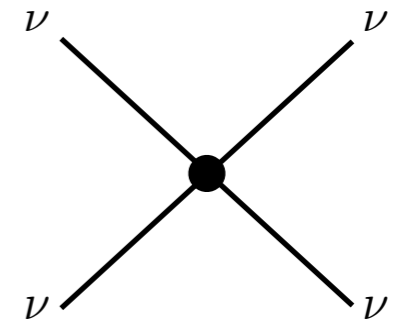


Cosmology of Slnu



$$\sigma \propto G_F^2$$

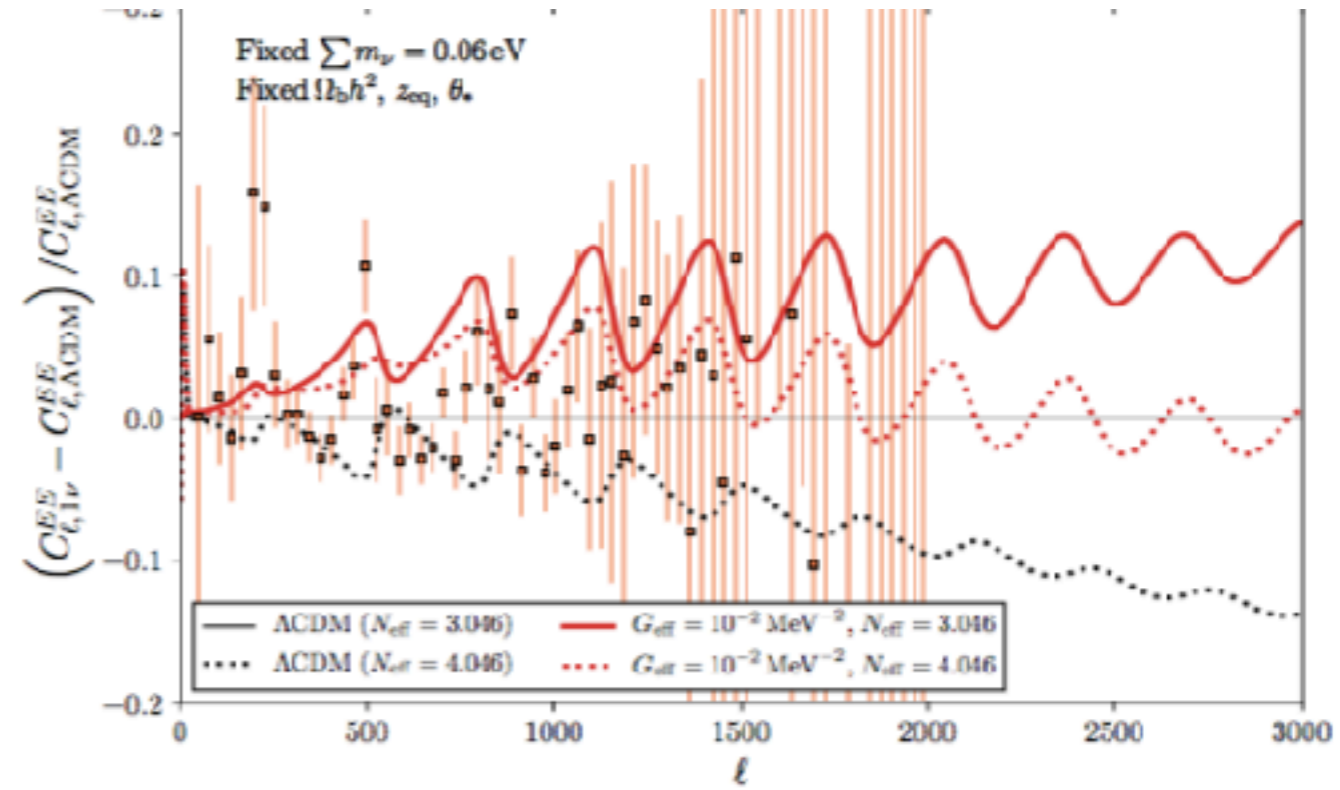
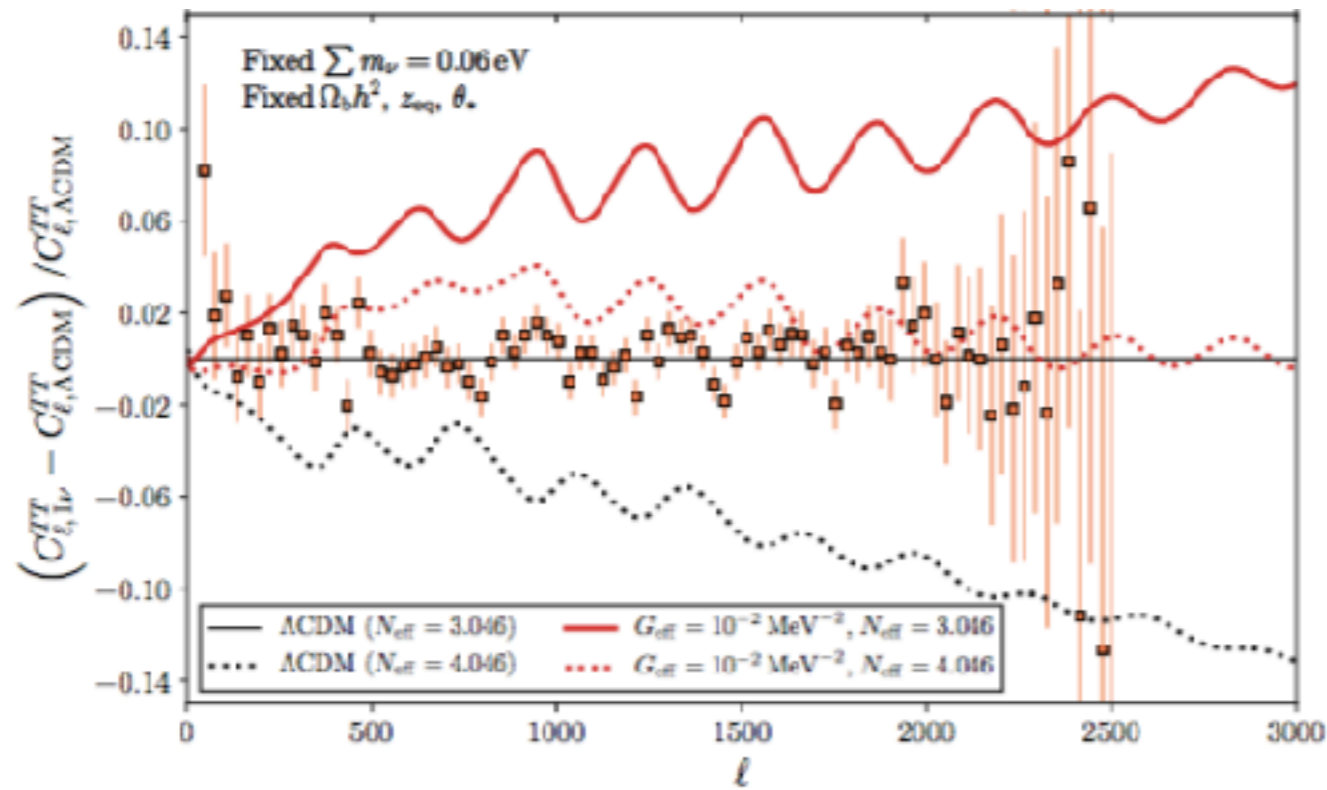
$$G_{\text{eff}} \simeq 10^9 G_F$$



$$\sigma \propto \frac{g_\phi^4}{M_\phi^4} \equiv G_{\text{eff}}^2$$

Cosmology of Slnu

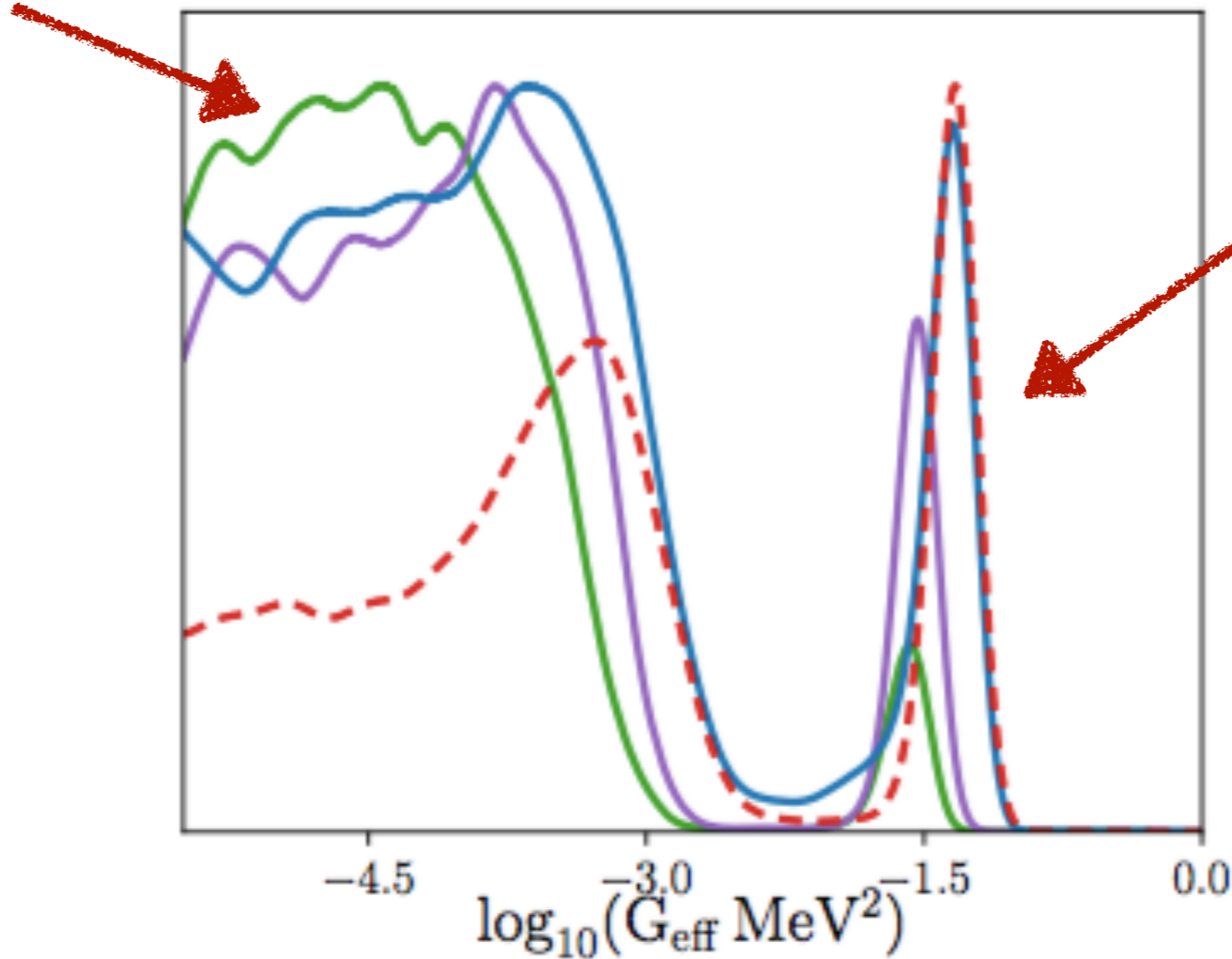
Gravitational potentials enhance photon perturbation \implies larger C_l



Kreisch et al 1902.00534

Cosmology of Slnu

Moderately Interacting
(MI)



Strongly Interacting
(SI)

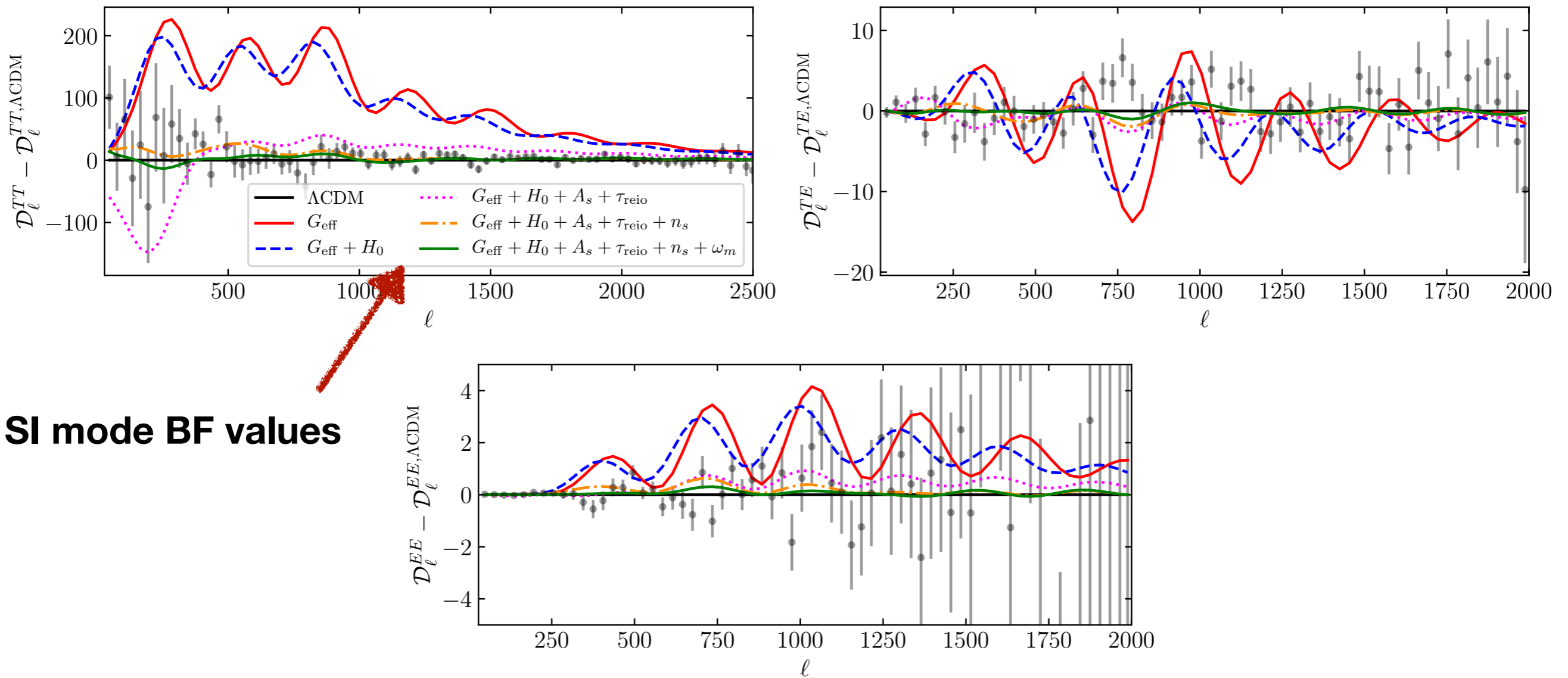
Kreisch et al 1902.00534

Questions:

- Origin of the SI mode?
- Why is there a valley in between?

Cosmology of Slnu: G_{eff} degeneracy w/ other params

P18 : TTTEEE + lowE | 3c + 0f



SI mode BF values

SI mode changes $l > 200$

SInu & Hubble

Phase shift $\rightarrow \phi_\nu \simeq 0.19\pi R_\nu \leftarrow$ Free-streaming neutrino

$$\ell \approx (m\pi - \phi_\nu) \frac{D_A^*}{r_s^*}$$

$$D_A^* = \int_0^{z^*} \frac{1}{H(z)} dz, \quad r_s^* = \int_{z^*}^{\infty} \frac{c_s(z)}{H(z)} dz$$

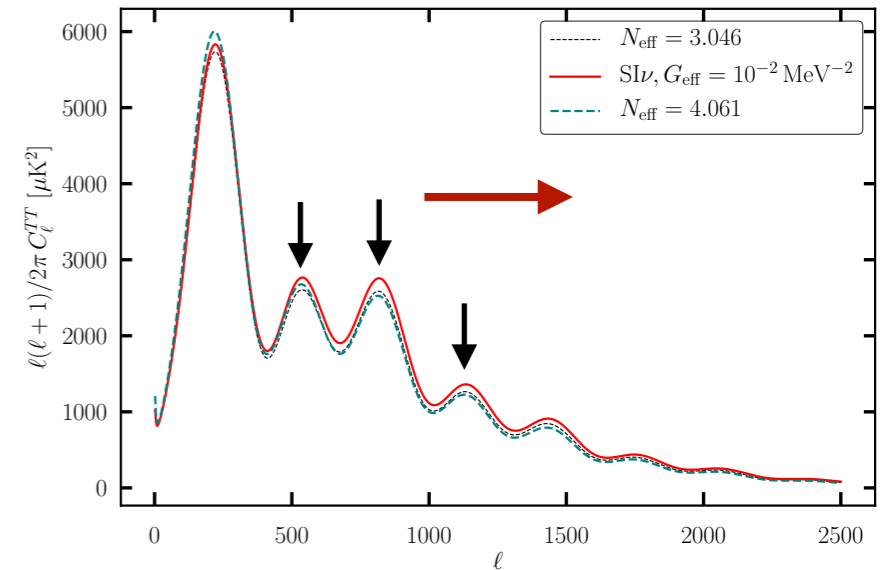
SInu phase shift increases ℓ



To keep the spectrum fixed, $\theta_* = \frac{r_s^*}{D_A^*}$ needs to increase



Larger H_0 decreases D_A^* & increases θ_*

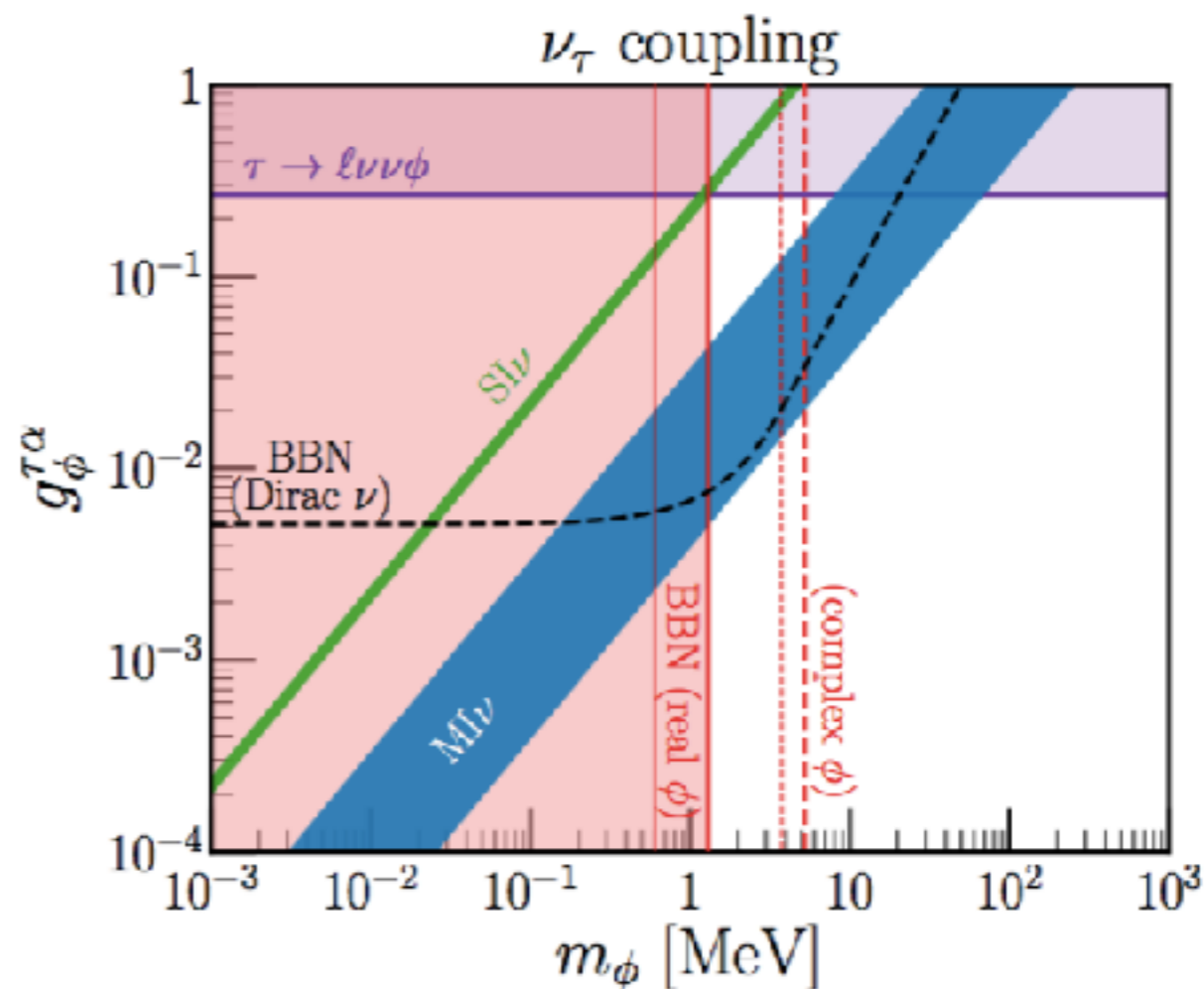
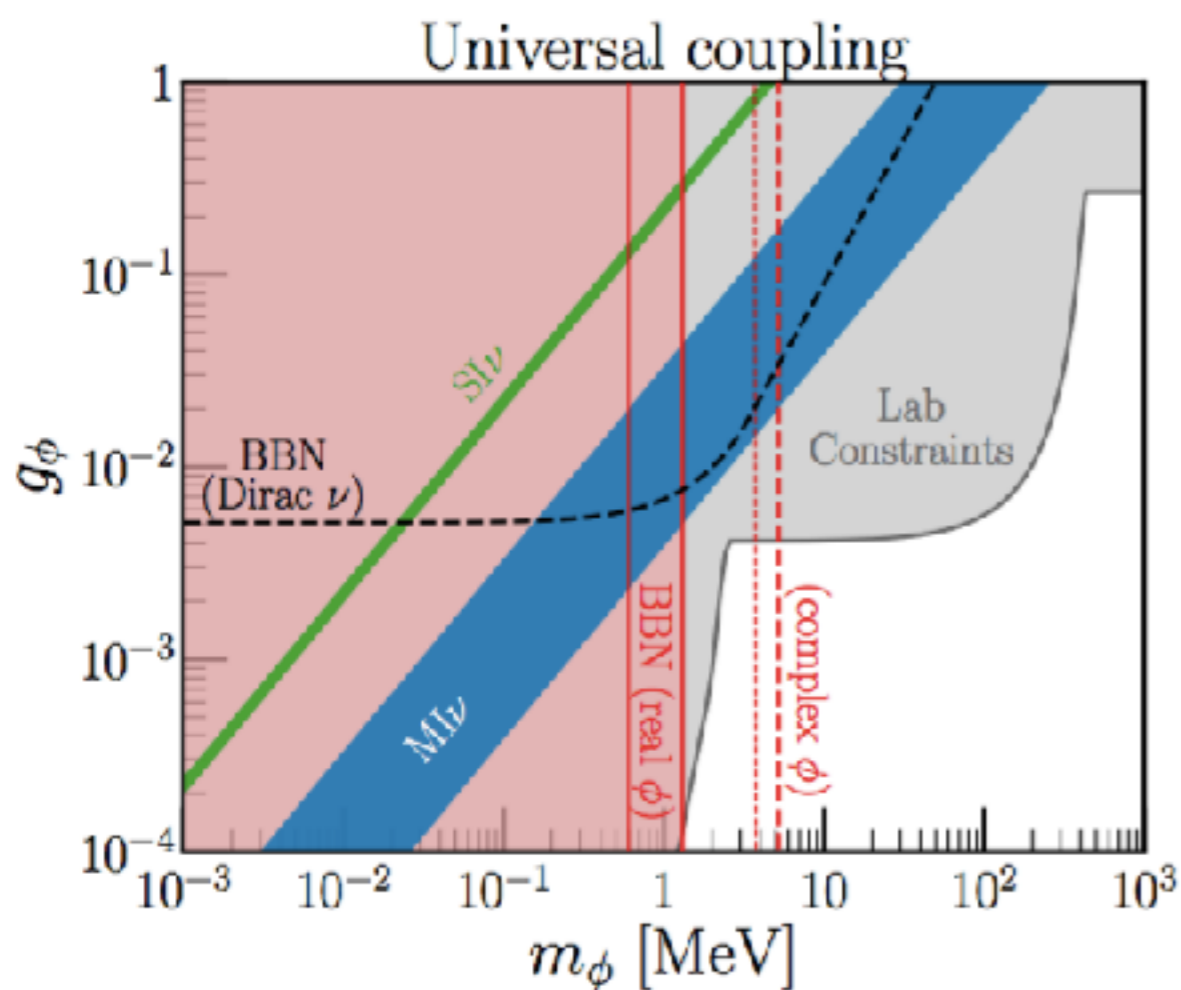


Laboratory Constraints

Example: ν self-interaction could come from a scalar (1905.02727)

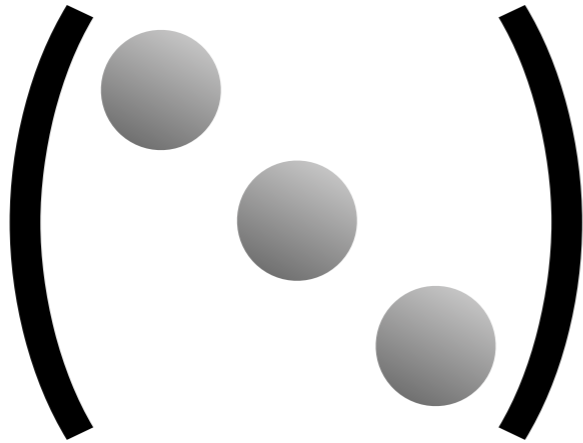
$$\mathcal{L} \supset -\frac{1}{2}m_\phi^2\phi^2 + \frac{1}{2}(g_\phi^{\alpha\beta}\nu_\alpha\nu_\beta\phi + \text{h.c.})$$

$$G_{\text{eff}} \equiv \frac{g_\phi^2}{m_\phi^2} = (10 \text{ MeV})^{-2} \left(\frac{g_\phi}{10^{-1}}\right)^2 \left(\frac{\text{MeV}}{m_\phi}\right)^2$$

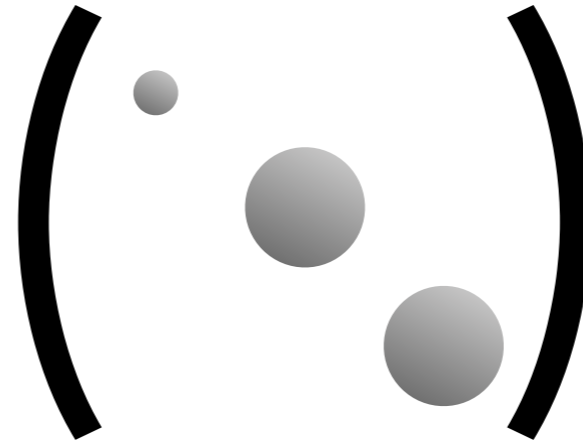


Flavor-specific Slnu

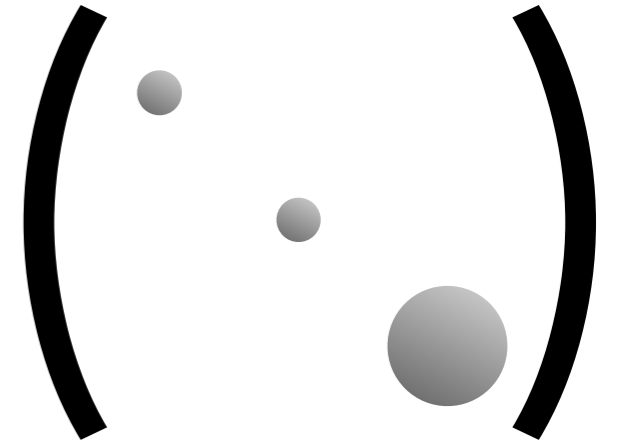
Flavor
Space



$3c+0f$



$2c+1f$

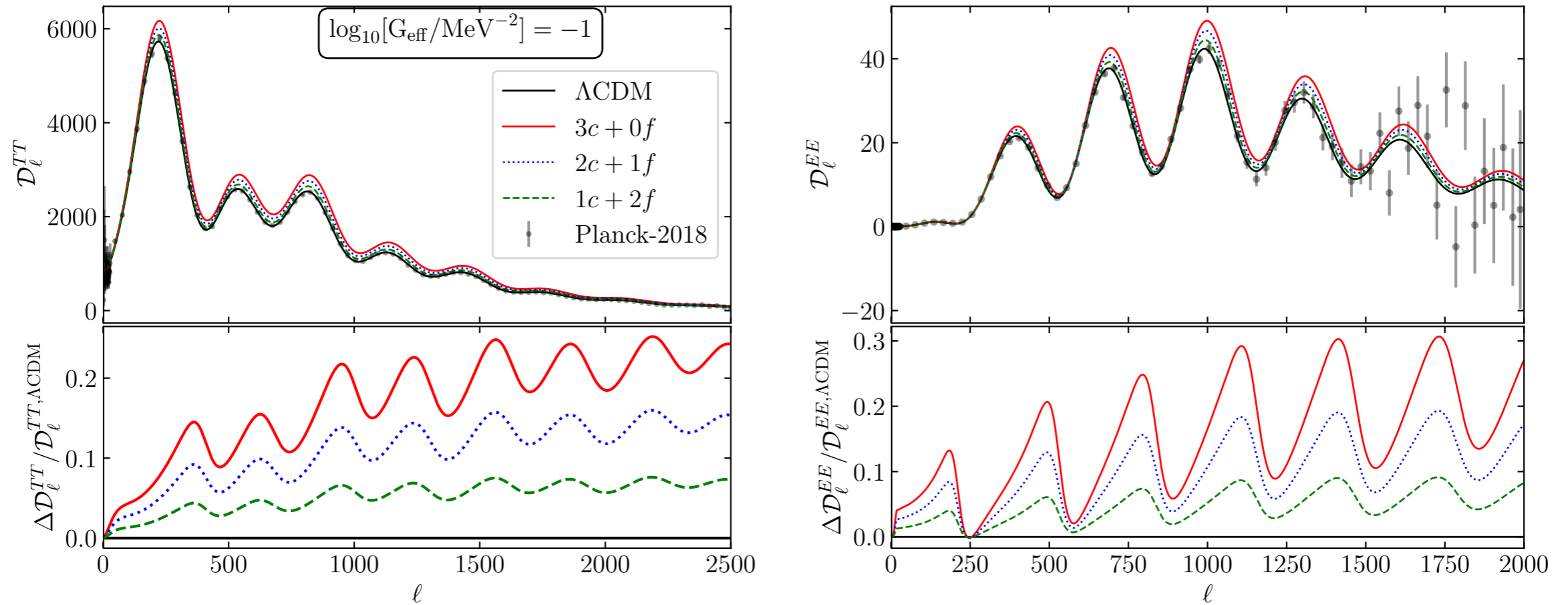


$1c+2f$

$$\mathcal{L} \supset G_{\text{eff}}^{(ijkl)} \bar{\nu}_i \nu_j \bar{\nu}_k \nu_l,$$

$$G_{\text{eff}}^{(ijkl)} \equiv \frac{g_{ij} g_{kl}}{M_\phi^2}$$

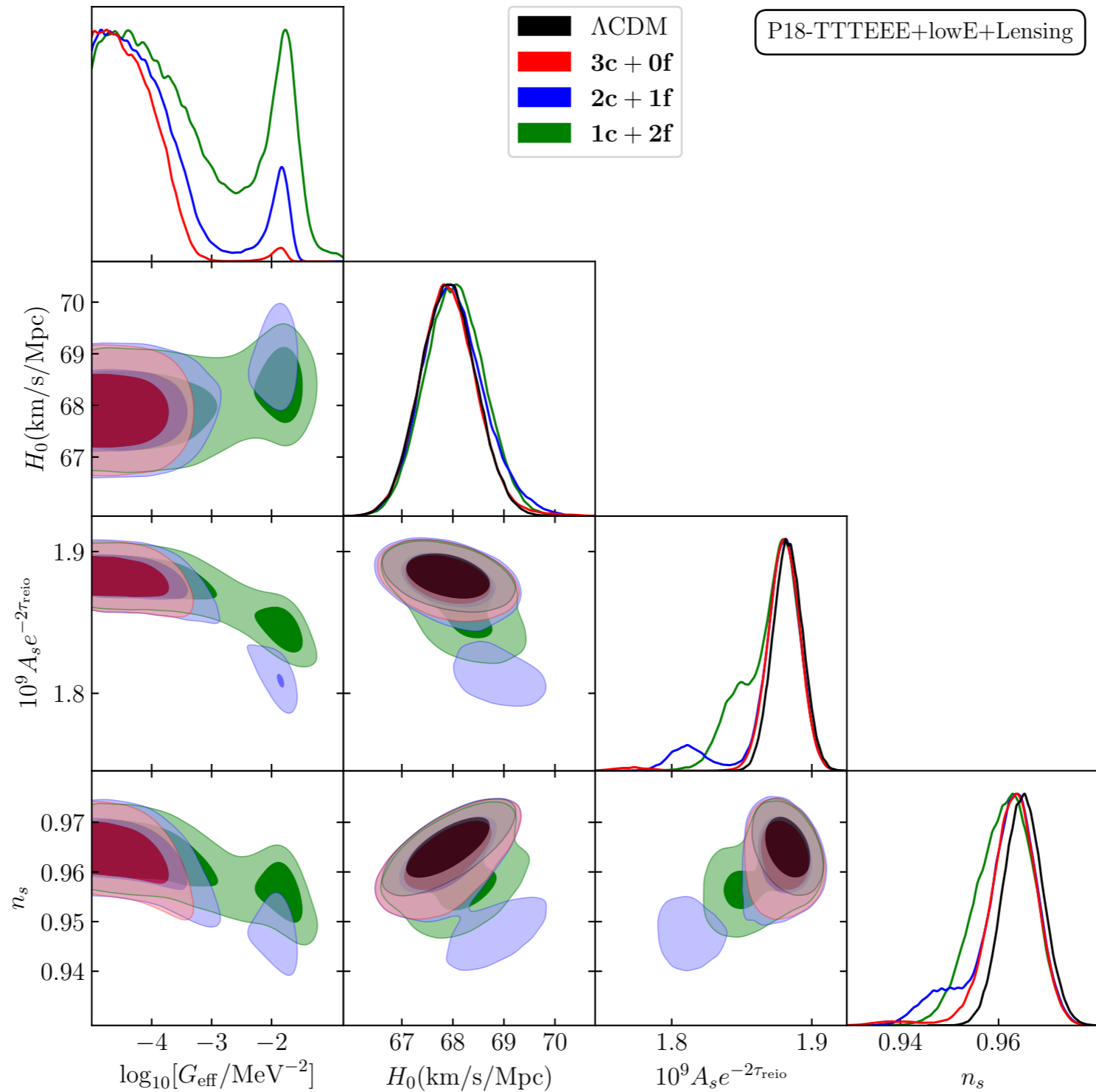
Flavor-specific S_{ν}



2011.12315

**Less number of interacting neutrinos reduces the change
in the CMB power spectra**

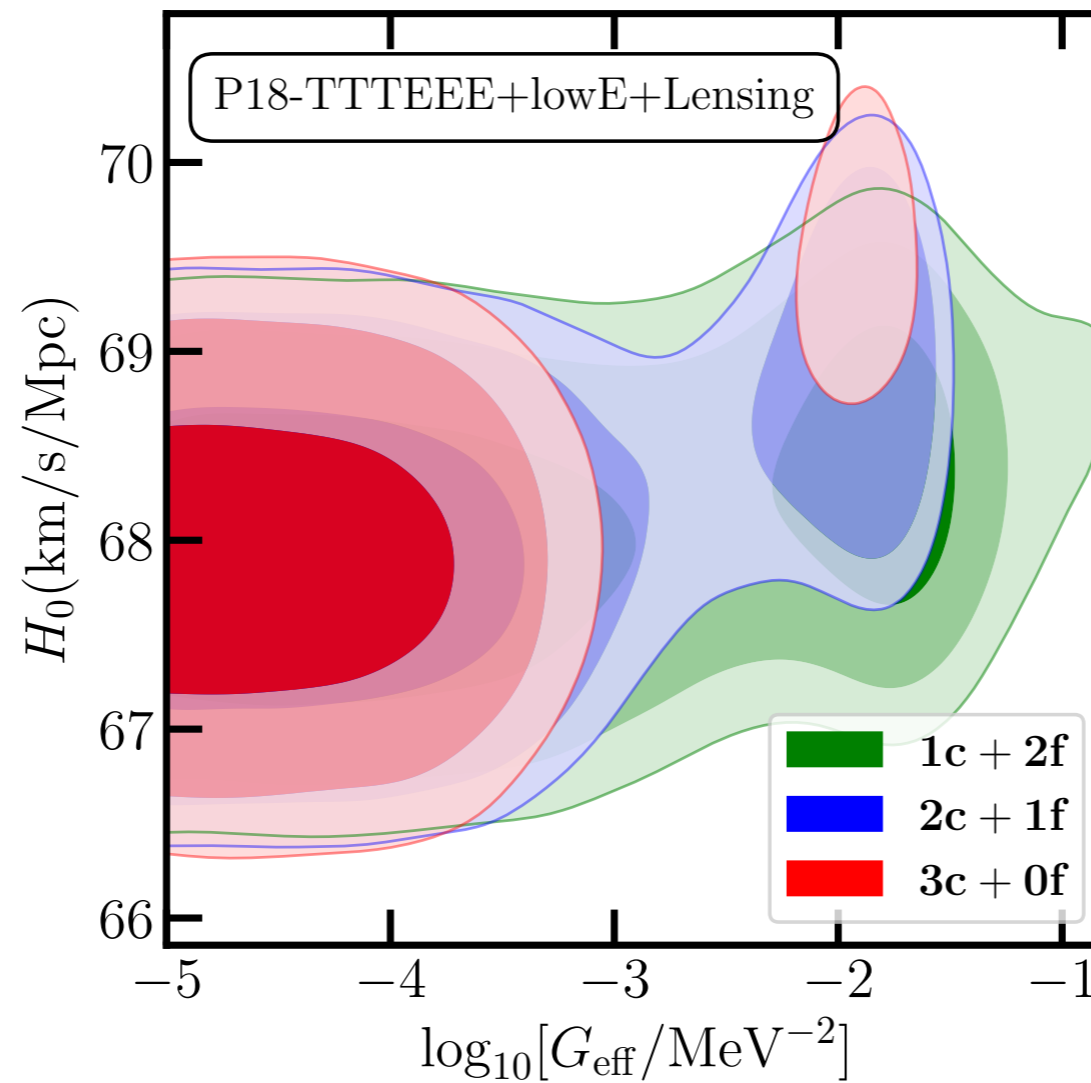
Flavor-specific Slnu



2011.12315

Flavor-specific S_{Inu} : effect on H_0

G_{eff} has positive correlation with H_0



$H_0 = 69.5 \pm 0.6$ km/s/Mpc for TTTEEE+lowE+lens (2011.12315)

$H_0 = 69.6 \pm 2.5$ (TRGB, 2002.01550)

$H_0 = 73.2 \pm 1.3$ (Riess et al, 2012.08534)

Conclusions & Future Outlook

- Neutrino self-interaction is phenomenologically motivated
- Flavor-universal large G_{eff} ($\sim 10^{-2} \text{ MeV}^{-2}$) is incompatible with BBN & laboratory data. However, flavor-specific interaction may still be allowed.
- The SI mode bestfit coupling in the flavor-specific case remains almost the same
- Changes in CMB spectrum is less with less #interacting nu, hence, more room to exploit the deg. with $A_s, n_s, \tau_{\text{reio}}, H_0 \implies$ **enhances the SI mode**
- More work needed with massive neutrinos and off-diagonal couplings

Parameter values

Table 4: Parameter values and 68% confidence limits in 2c + 1f.

| Parameters | TT+lowE | | TTTEEE+lowE | |
|---|----------------------|---------------------|---------------------|----------------------|
| | SI | MI | SI | MI |
| $\Omega_b h^2$ | 0.022 ± 0.00027 | 0.022 ± 0.00021 | 0.022 ± 0.00016 | 0.022 ± 0.00015 |
| $\Omega_c h^2$ | 0.1211 ± 0.0023 | 0.1203 ± 0.002 | 0.1205 ± 0.0014 | 0.1201 ± 0.0013 |
| $100\theta_s$ | 1.0452 ± 0.00059 | 1.0419 ± 0.0005 | 1.045 ± 0.00076 | 1.0419 ± 0.00031 |
| $\ln(10^{10} A_s)$ | 2.99 ± 0.0179 | 3.036 ± 0.01714 | 3 ± 0.0167 | 3.042 ± 0.0161 |
| n_s | 0.9407 ± 0.0079 | 0.9596 ± 0.0068 | 0.9473 ± 0.0046 | 0.9628 ± 0.005 |
| τ_{reio} | 0.0501 ± 0.008 | 0.0516 ± 0.0079 | 0.0538 ± 0.0077 | 0.0538 ± 0.0077 |
| $\log_{10}(G_{\text{eff}}/\text{MeV}^{-2})$ | -1.69 ± 0.2 | -4.03 ± 0.6 | -1.93 ± 0.24 | -4.24 ± 0.5 |
| $H_0(\text{ km s}^{-1}\text{Mpc}^{-1})$ | 68.34 ± 1.00 | 67.57 ± 0.92 | 68.81 ± 0.63 | 67.83 ± 0.6 |
| σ_8 | 0.823 ± 0.01 | 0.824 ± 0.009 | 0.829 ± 0.0079 | 0.824 ± 0.0075 |

Table 5: Parameter values and 68% confidence limits in 1c + 2f.

| Parameters | TT+lowE | | TTTEEE+lowE | |
|---|----------------------|---------------------|---------------------|---------------------|
| | SI | MI | SI | MI |
| $\Omega_b h^2$ | 0.022 ± 0.00023 | 0.022 ± 0.00021 | 0.022 ± 0.00015 | 0.022 ± 0.00015 |
| $\Omega_c h^2$ | 0.1207 ± 0.0021 | 0.1203 ± 0.002 | 0.1203 ± 0.0014 | 0.1201 ± 0.0013 |
| $100\theta_s$ | 1.0434 ± 0.00062 | 1.0419 ± 0.0004 | 1.043 ± 0.00058 | 1.0419 ± 0.0003 |
| $\ln(10^{10} A_s)$ | 3.01 ± 0.0179 | 3.037 ± 0.01664 | 3.024 ± 0.0166 | 3.042 ± 0.016 |
| n_s | 0.9513 ± 0.0069 | 0.9609 ± 0.0059 | 0.9553 ± 0.0049 | 0.963 ± 0.005 |
| τ_{reio} | 0.051 ± 0.008 | 0.0519 ± 0.008 | 0.0539 ± 0.0076 | 0.0539 ± 0.0077 |
| $\log_{10}(G_{\text{eff}}/\text{MeV}^{-2})$ | -1.75 ± 0.4 | -3.94 ± 0.6 | -1.9 ± 0.37 | -4.06 ± 0.6 |
| $H_0(\text{ km s}^{-1}\text{Mpc}^{-1})$ | 67.9 ± 1.00 | 67.56 ± 0.93 | 68.3 ± 0.62 | 67.83 ± 0.61 |
| σ_8 | 0.821 ± 0.01 | 0.823 ± 0.009 | 0.825 ± 0.0083 | 0.824 ± 0.0075 |

Parameter values

Table 11: Parameter values and 68% confidence limits in **2c + 1f**.

| Parameters | TTTEEE+lowE+lens | | TTTEEE+lowE+lens+BAO+ H_0 | |
|---|---------------------|---------------------|-----------------------------|---------------------|
| | SI | MI | SI | MI |
| $\Omega_b h^2$ | 0.022 ± 0.0001 | 0.022 ± 0.00014 | 0.022 ± 0.0001 | 0.022 ± 0.00013 |
| $\Omega_c h^2$ | 0.1202 ± 0.0013 | 0.1199 ± 0.0012 | 0.12 ± 0.001 | 0.1188 ± 0.0009 |
| $100\theta_s$ | 1.045 ± 0.0008 | 1.0419 ± 0.0003 | 1.045 ± 0.00068 | 1.042 ± 0.00029 |
| $\ln(10^{10} A_s)$ | 3 ± 0.0158 | 3.041 ± 0.014 | 3 ± 0.0151 | 3.044 ± 0.0145 |
| n_s | 0.9476 ± 0.0043 | 0.9629 ± 0.0048 | 0.9483 ± 0.004 | 0.966 ± 0.0046 |
| τ_{reio} | 0.0541 ± 0.0074 | 0.0536 ± 0.0072 | 0.0544 ± 0.007 | 0.0565 ± 0.0071 |
| $\log_{10}(G_{\text{eff}}/\text{MeV}^{-2})$ | -1.96 ± 0.26 | -4.22 ± 0.51 | -1.91 ± 0.22 | -4.22 ± 0.52 |
| $H_0(\text{km s}^{-1}\text{Mpc}^{-1})$ | 68.87 ± 0.58 | 67.9 ± 0.53 | 69.08 ± 0.42 | 68.47 ± 0.4 |
| σ_8 | 0.829 ± 0.007 | 0.823 ± 0.006 | 0.827 ± 0.0065 | 0.821 ± 0.0059 |

Table 12: Parameter values and 68% confidence limits in **1c + 2f**.

| Parameters | TTTEEE+lowE+lens | | TTTEEE+lowE+lens+BAO+ H_0 | |
|---|---------------------|---------------------|-----------------------------|---------------------|
| | SI | MI | SI | MI |
| $\Omega_b h^2$ | 0.022 ± 0.0001 | 0.022 ± 0.00014 | 0.022 ± 0.0001 | 0.022 ± 0.00013 |
| $\Omega_c h^2$ | 0.1201 ± 0.0012 | 0.1199 ± 0.0012 | 0.12 ± 0.0009 | 0.1188 ± 0.0009 |
| $100\theta_s$ | 1.043 ± 0.0006 | 1.0419 ± 0.0003 | 1.043 ± 0.00056 | 1.042 ± 0.00029 |
| $\ln(10^{10} A_s)$ | 3.023 ± 0.0153 | 3.041 ± 0.015 | 3 ± 0.0151 | 3.045 ± 0.0142 |
| n_s | 0.9555 ± 0.0046 | 0.9633 ± 0.0045 | 0.9572 ± 0.004 | 0.966 ± 0.0042 |
| τ_{reio} | 0.0536 ± 0.0073 | 0.0536 ± 0.0074 | 0.0554 ± 0.007 | 0.0566 ± 0.0071 |
| $\log_{10}(G_{\text{eff}}/\text{MeV}^{-2})$ | -1.91 ± 0.37 | -4.04 ± 0.61 | -1.86 ± 0.36 | -4.03 ± 0.61 |
| $H_0(\text{km s}^{-1}\text{Mpc}^{-1})$ | 68.35 ± 0.56 | 67.9 ± 0.54 | 68.75 ± 0.41 | 68.48 ± 0.41 |
| σ_8 | 0.824 ± 0.007 | 0.823 ± 0.006 | 0.822 ± 0.0071 | 0.821 ± 0.006 |