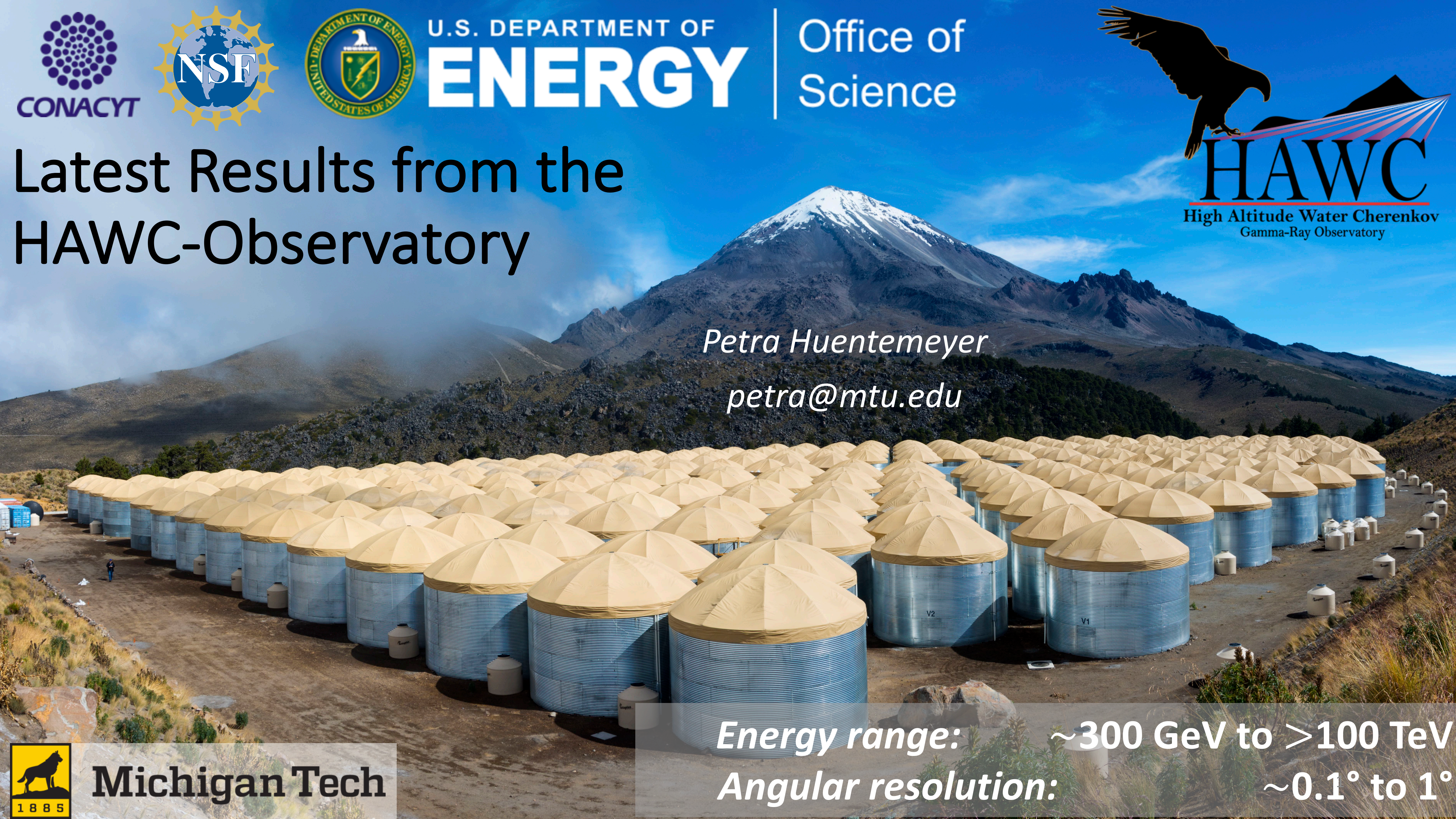


CONACYT

NSF

U.S. DEPARTMENT OF ENERGY | Office of Science

HAWC
High Altitude Water Cherenkov
Gamma-Ray Observatory

Latest Results from the HAWC-Observatory

Petra Huentemeyer

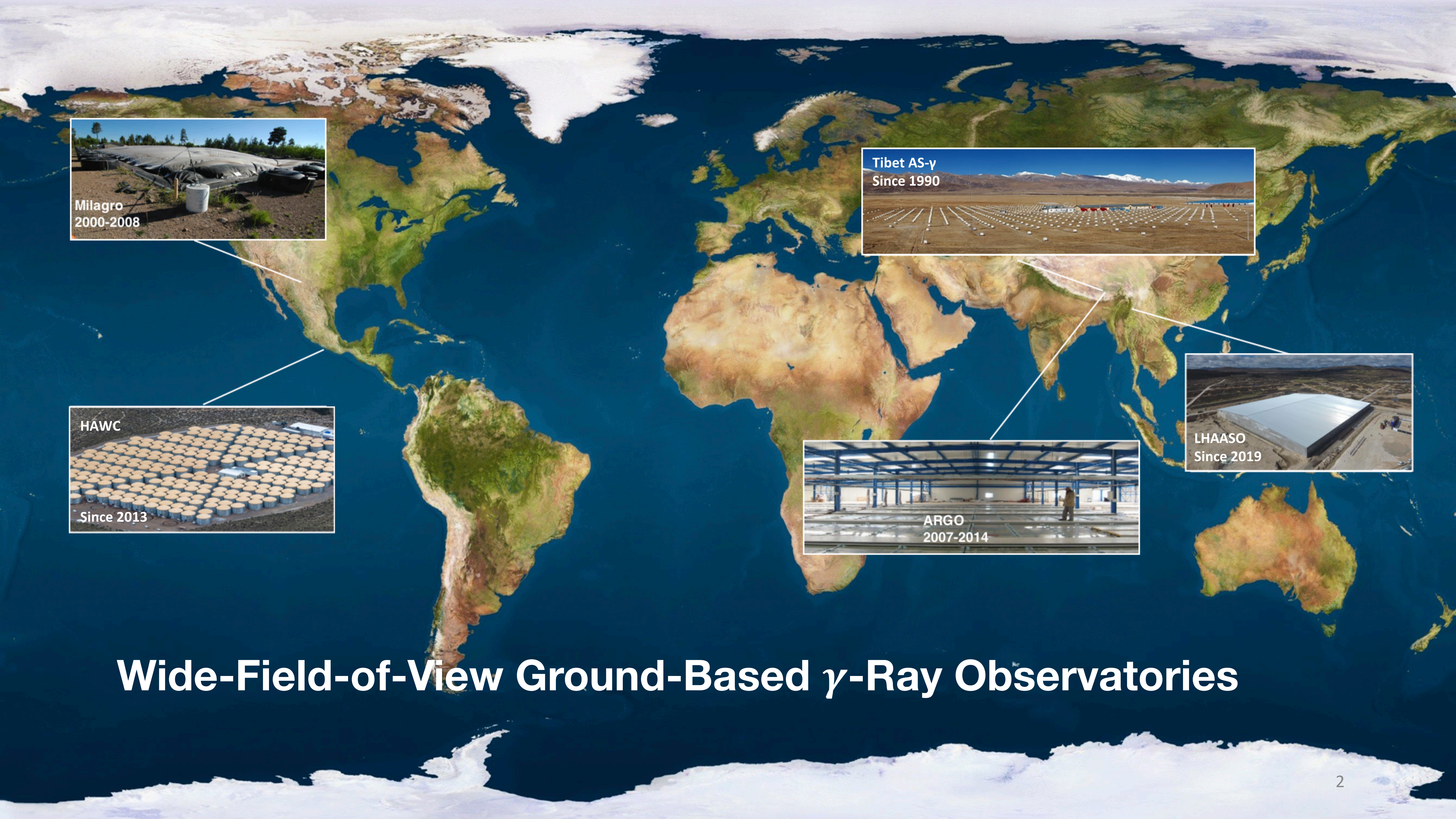
petra@mtu.edu

Energy range: ~300 GeV to >100 TeV

Angular resolution: ~0.1° to 1°

Michigan Tech

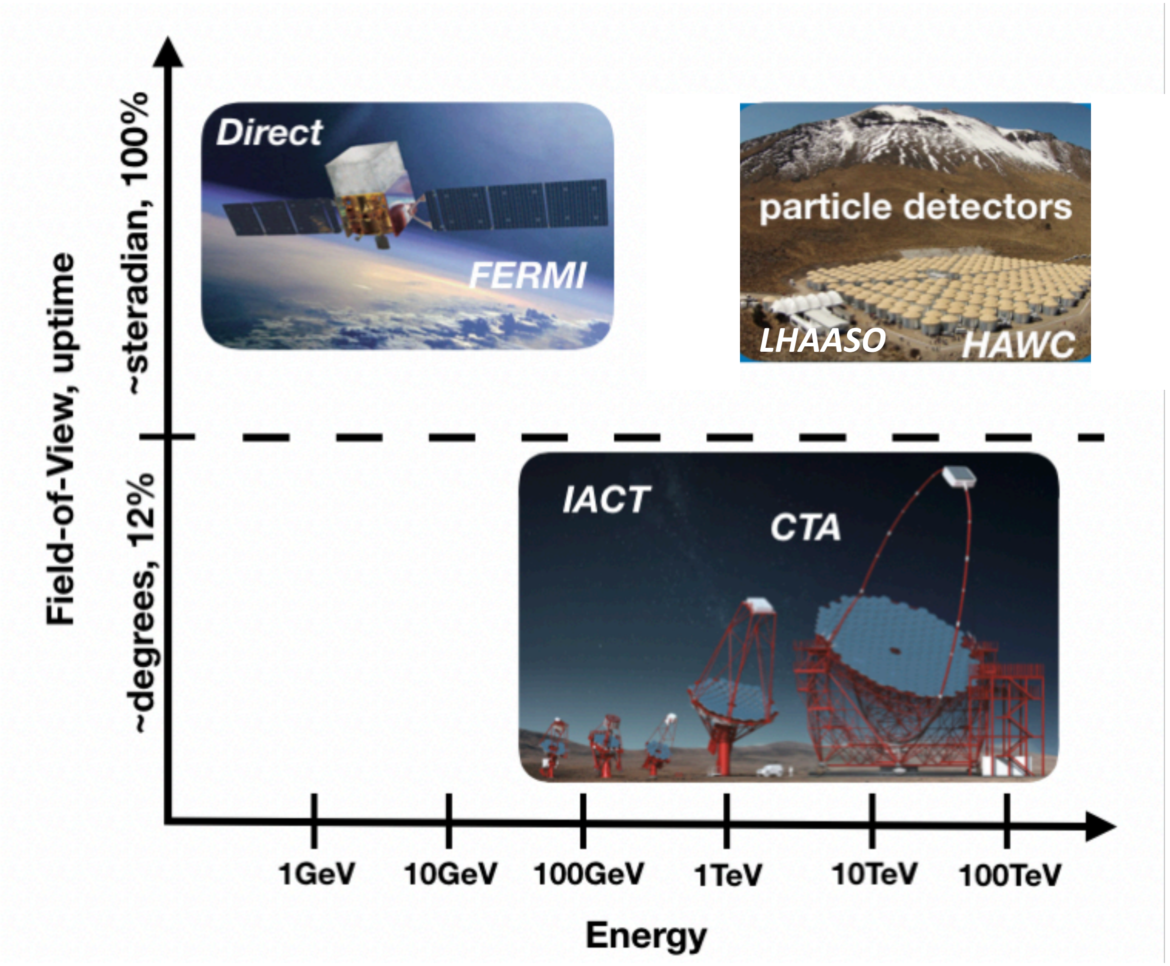
Wide-Field-of-View Ground-Based γ-Ray Observatories



Duty Cycle, Field-of-View & Energy

3 Main Features

- High duty cycle (> 95% uptime)
  - ✓ Transients
- Wide field-of-view
  - ✓ Extended and large scale emission
- Good Sensitivity, Angular & Energy Resolution > 10 TeV
  - ✓ Highest energy accelerators

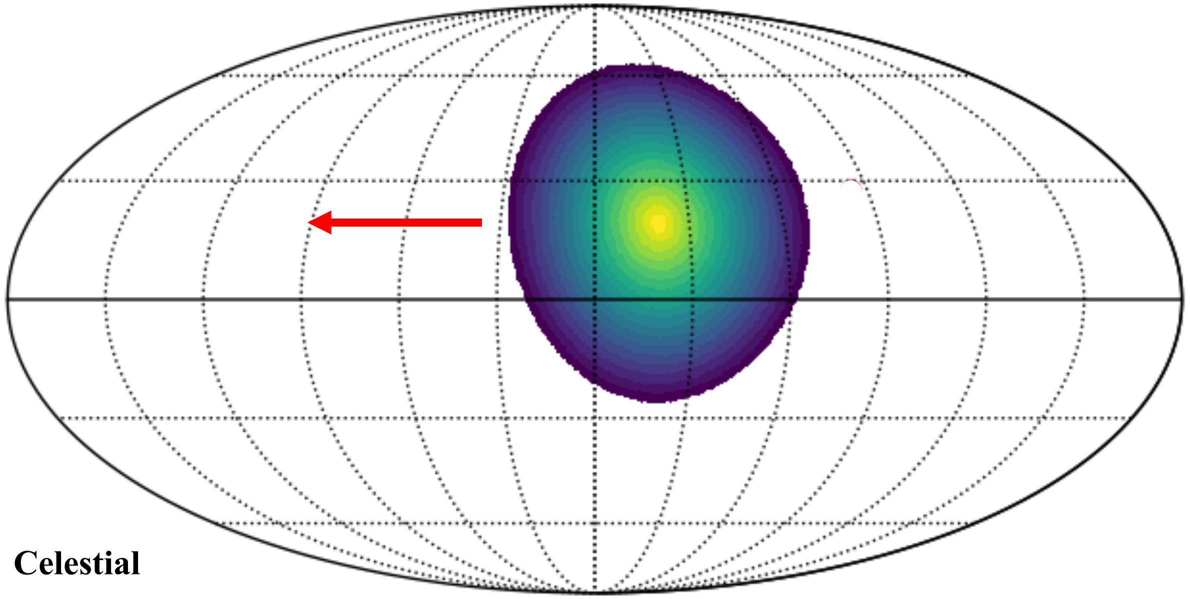


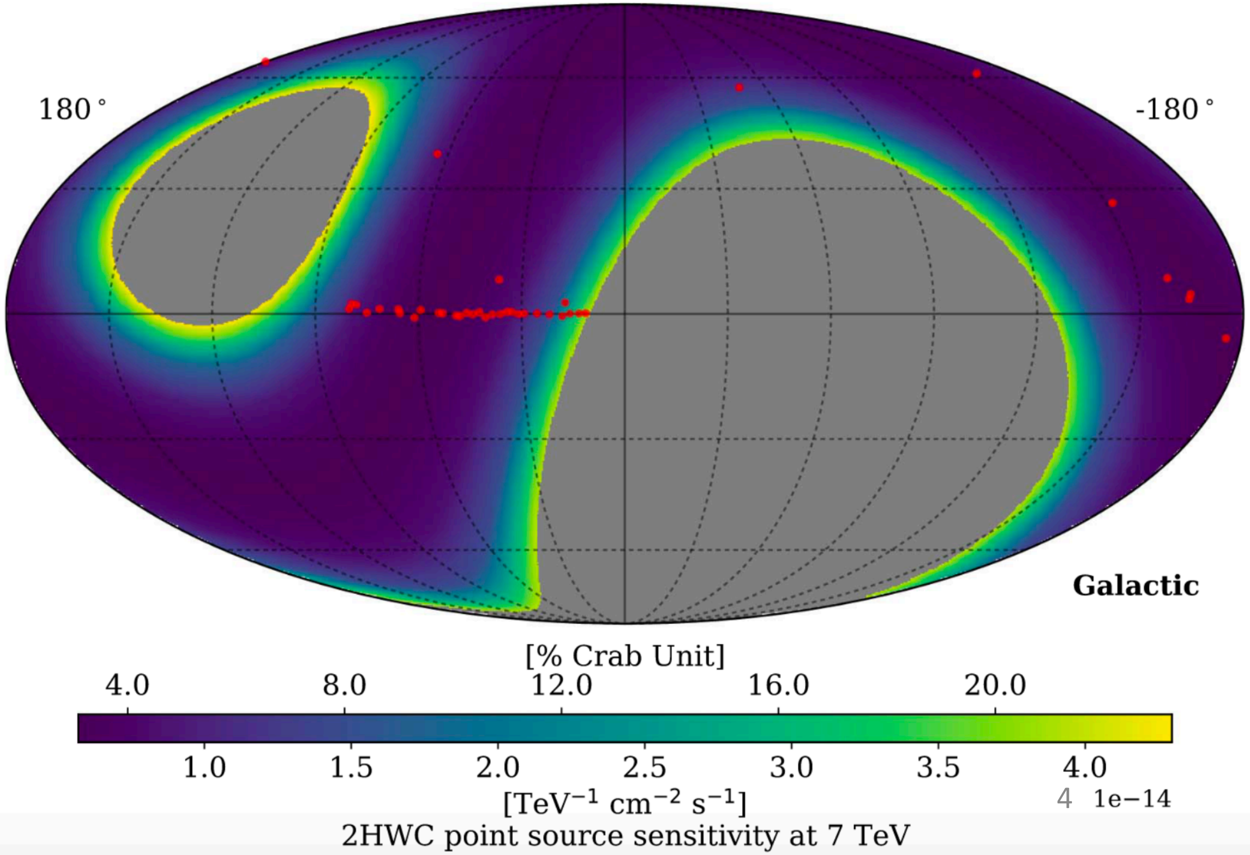
Wide Field-Of-View

- HAWC (almost) continuously observes the sky as it transits over its zenith
- Instantaneous field of view of HAWC~ 1.8 sr (~15% of the sky)

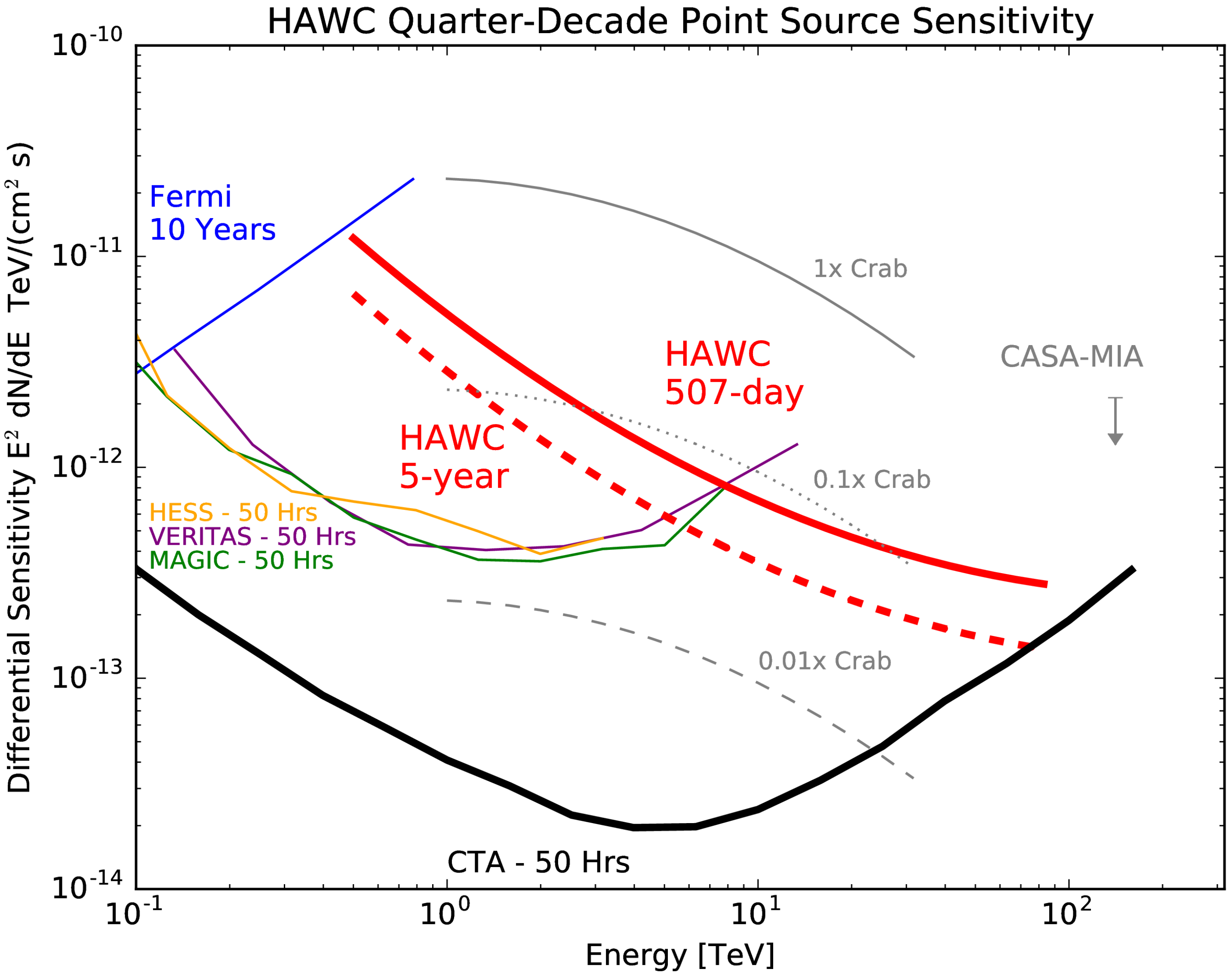
  → surveys ~ 8.4 sr / day (~2/3 of the sky).

ApJ, Vol. 843, Is. 1, id. 40, 21pp. (2017)





Sensitivity vs Energy

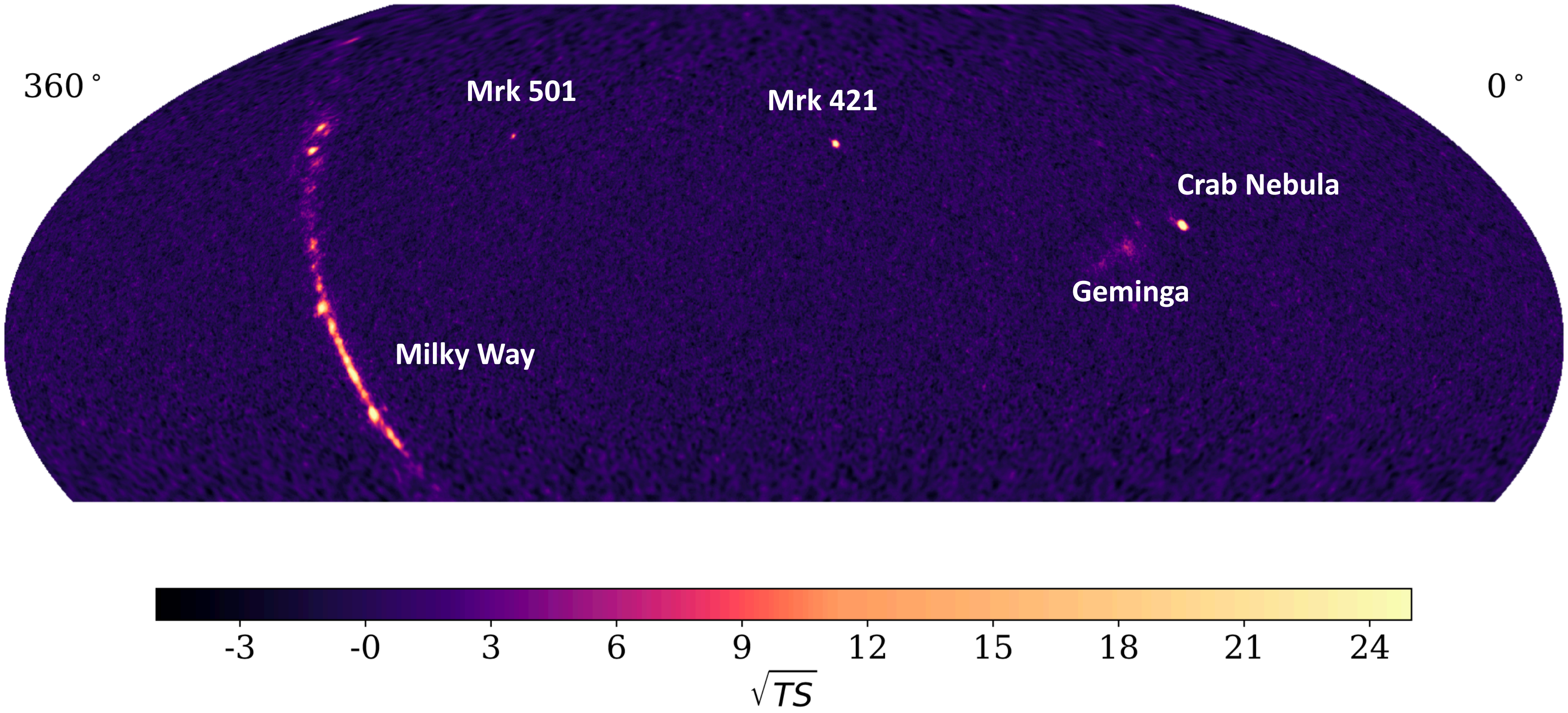


ApJ, Vol. 843, Is. 1, id. 39, 17 pp. (2017)

Main Features → Latest HAWC Results

- Survey capabilities
  - ✓ 3rd HAWC catalog
  - ✓ Unexpected emission sources
- Extended and large scale emission sensitivity
  - ✓ New source class: halos
  - ✓ Molecular Clouds
- High-energy γ-ray sensitivity
  - ✓ Several PeVatron candidates
  - ✓ Test of fundamental physics in unique phase space

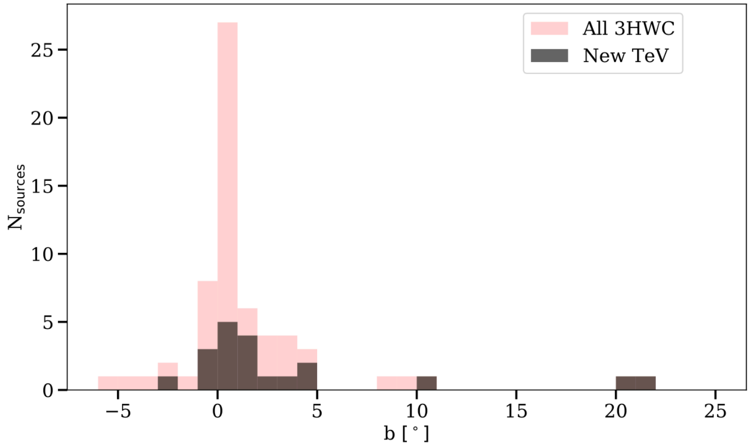
Survey of the Northern Sky: 3HWC Catalog

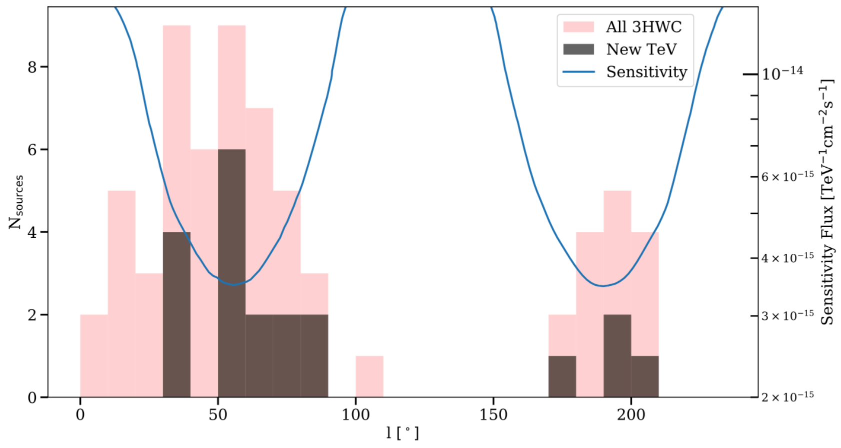


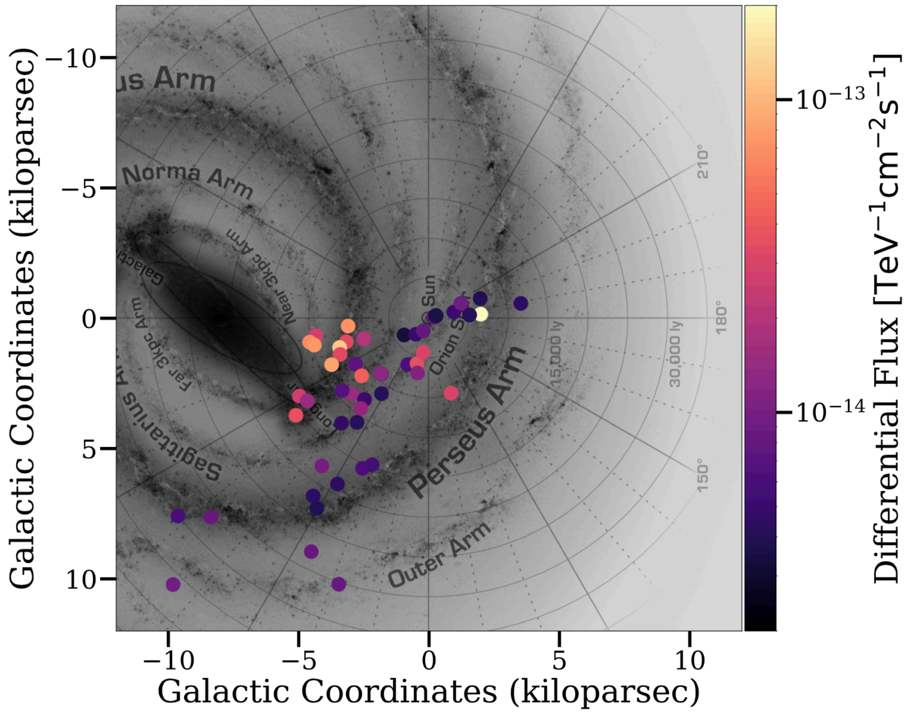
ApJ, Vol. 905, Is. 1, id.76, 14pp. (2020)

- Catalog from 1523 days of data: most sensitive survey of the northern γ-ray sky > several TeV
- 65 sources detected at > 5σ:
  - 20 sources > 1°away from previously detected TeV sources,
  - 14 of these have potential counterpart in the 4th Fermi-LAT catalog

3HWC Catalog: Spatial Distribution of Sources



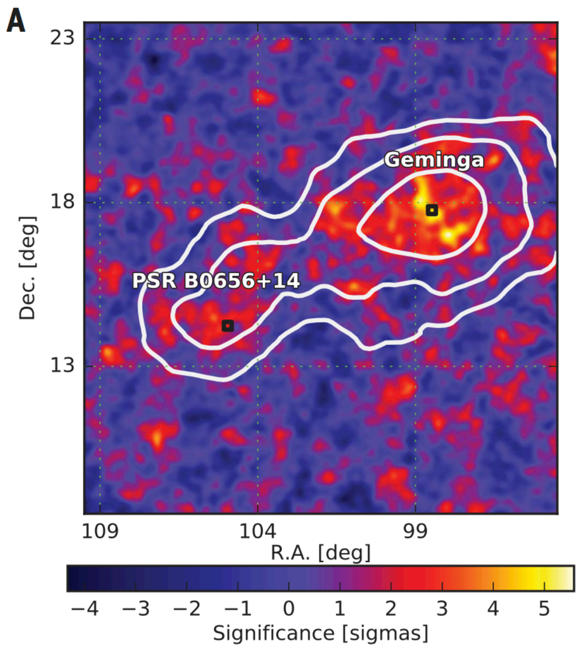




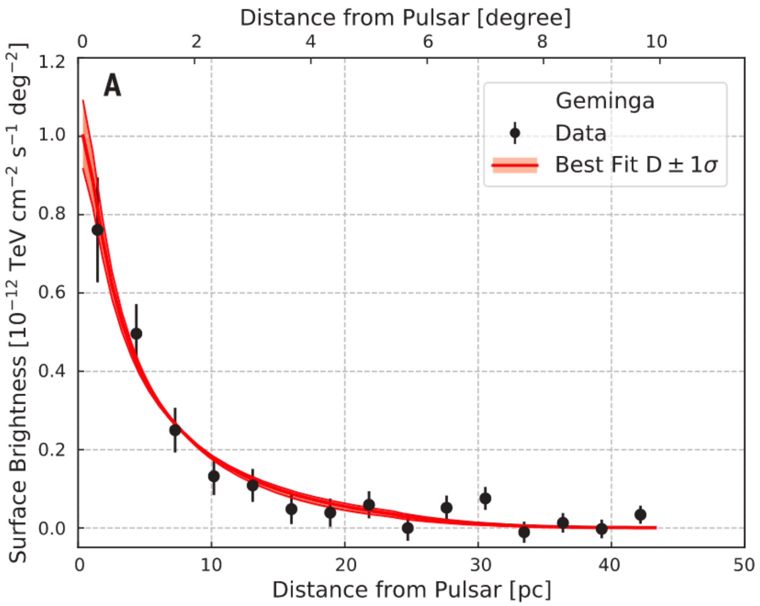
ApJ, Vol. 905, Is. 1, id.76, 14 pp. (2020)

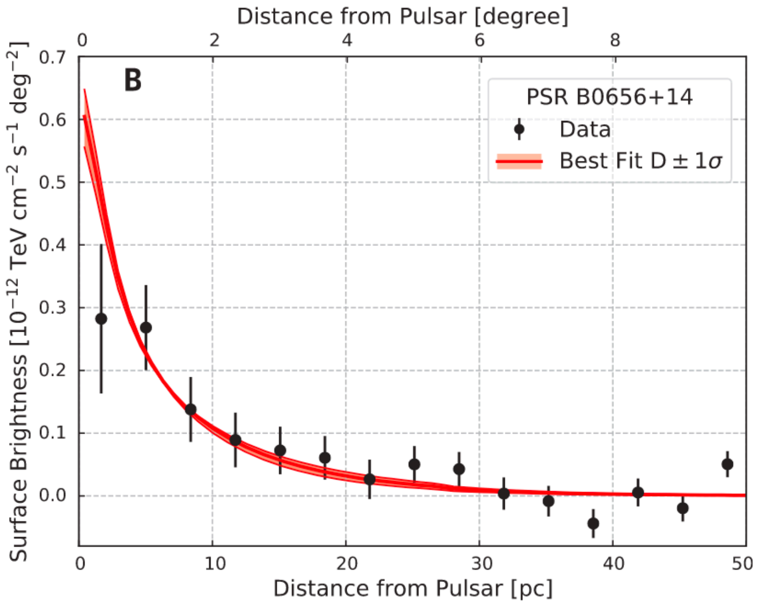
- Majority along the Galactic plane
- Due to its location, HAWC most sensitive towards the Galactic anti-center region and, to the inner Galaxy
- Significant fraction of 3HWC sources found near pulsars listed in the ATNF catalog (Two of these, PSR J0631+1036 and PSR J1740+1000, have not previously been connected with TeV emission)

New Source Class: Halos



[Science 358, 911–914 (2017)]



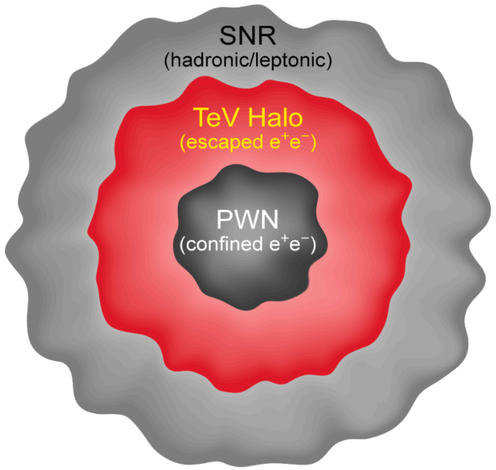


- In 2017, HAWC reported the detection of extended TeV γ-ray emission coincident with the locations of two nearby middle-aged pulsars: Geminga and PSR B0656+14 (inside Monogem ring) [[Science 358, 911–914 (2017)]]
- Observations demonstrated that these pulsars are indeed local sources of accelerated leptons, and the surrounding emission profile can be used to constrain the diffusion of particles away from their sources → much slower than previously assumed

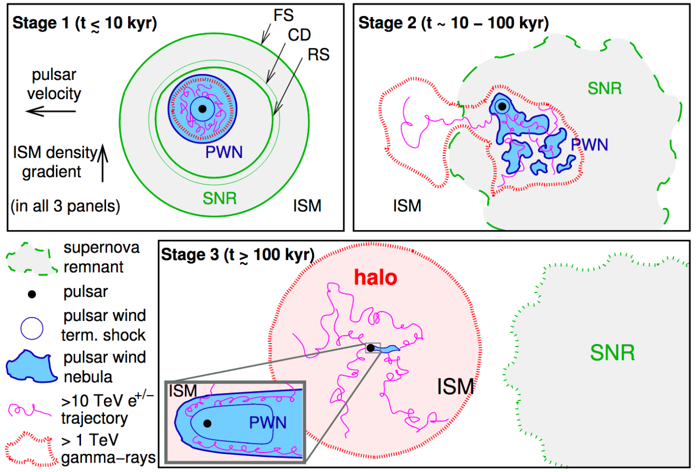
New Source Class: Halos

The Geminga halo discovery and the discovery of several extended TeV PWNe by H.E.S.S. (A&A 612, A2 (2018)), lead to the hypothesis that extended ***"Halos"*** are a common feature of pulsars [PRD 96, 103016 (2017); PRL 120, 121101 (2018); PRD, 100, 043016 (2019); Astro2020; BAAS, Vol. 51, Is. 3, id. 311 (2019) ; A&A 636, A113 (2020)]

PRD, 100, 043016 (2019)



A&A 636, A113 (2020)



- ***Interpretations:***
  - Observed $\gamma-$ray emission due to IC up-scattering of CMB photons by relativistic e^-, e^+ that have escaped from the PWN, but remain trapped in a larger region where diffusion is ***inhibited*** compared to the interstellar medium
  - Only, form around very old pulsar (at least 100 kyr old) that either left their SNR shell or whose SNR shell already dissipated, allowing relativistic e^-, e^+ to diffuse freely in the vicinity of the pulsar
  - Distinct from (classical) PWNe, in that the $e^- - e^+$ plasma escaped from the x-ray PWN.
- Also detected at lower γ-ray energies by Fermi-LAT: ***γ-ray halos*** [Rev. D 100, 123015 (2019)]

Halos Candidates in the 3HWC Catalog

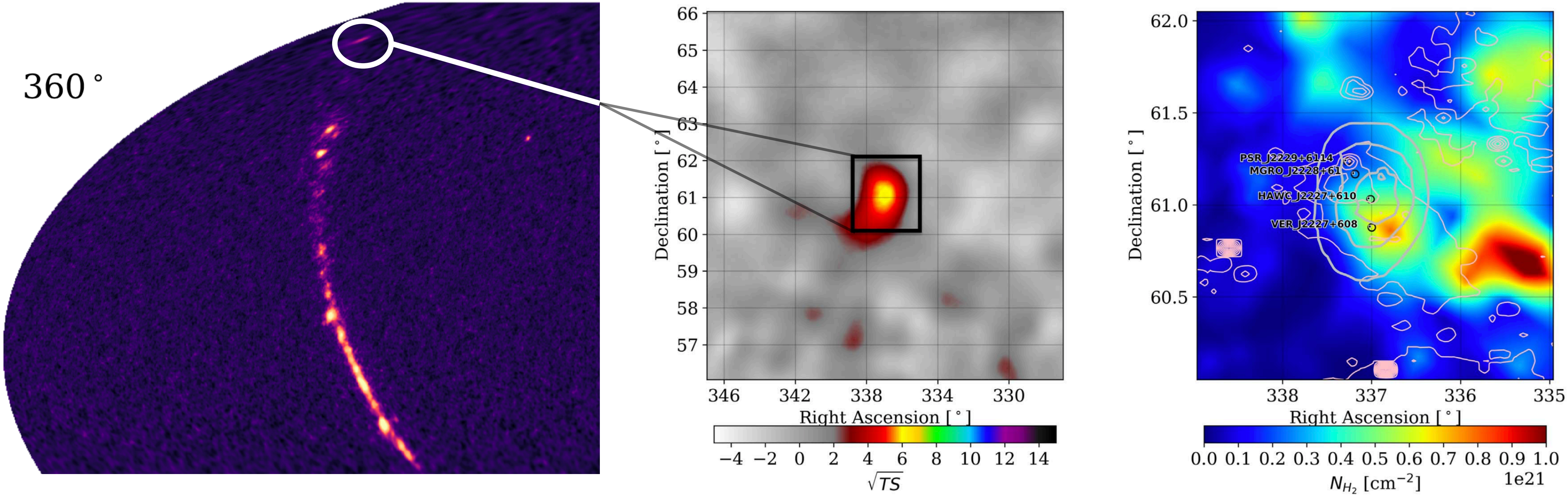
- Using similar criteria as PRD 96, 103016 (2017), a list of pulsars is created within 3HWC catalog that are likely candidates to have a TeV Halo: ATNF pulsars with
  - ages between 100kyr and 400kyr
  - declinations between $-25°$ and $+64°$
  - estimated spindown flux of at least 1% of that of the Geminga pulsar.

→ 16 such pulsars, 8 spatially coincident with at least one 3HWC source (within 1◦)

HAWC	l [°]	b [°]	Pulsar	Age [kyr]	$\dot{E}$ [erg s^{-1}]	Distance [kpc]	Separation [°]	TeVCat
3HWC J0540+228	184.58	-4.13	B0540+23	253.0	4.09e+34	1.56	0.83	HAWC J0543+233
3HWC J0543+231	184.67	-3.52	B0540+23	253.0	4.09e+34	1.56	0.36	HAWC J0543+233
3HWC J0631+169	195.63	3.45	J0633+1746	342.0	3.25e+34	0.19	0.95	Geminga
3HWC J0634+180	195.00	4.62	J0633+1746	342.0	3.25e+34	0.19	0.38	Geminga Pulsar
3HWC J0659+147	200.60	8.40	B0656+14	111.0	3.8e+34	0.29	0.51	2HWC J0700+143
3HWC J0702+147	200.91	9.01	B0656+14	111.0	3.8e+34	0.29	0.77	2HWC J0700+143
3HWC J1739+099	33.89	20.34	J1740+1000	114.0	2.32e+35	1.23	0.13	...
3HWC J1831-095	22.13	0.02	J1831−0952	128.0	1.08e+36	3.68	0.27	HESS J1831-098
3HWC J1912+103	44.50	0.15	J1913+1011	169.0	2.87e+36	4.61	0.31	HESS J1912+101
3HWC J1923+169	51.58	0.89	J1925+1720	115.0	9.54e+35	5.06	0.67	...
3HWC J1928+178	52.93	0.20	J1925+1720	115.0	9.54e+35	5.06	0.85	2HWC J1928+177
3HWC J2031+415	80.21	1.14	J2032+4127	201.0	1.52e+35	1.33	0.11	TeV J2032+4130

ApJ, Vol. 905, Is. 1, id.76, 14 pp. (2020)

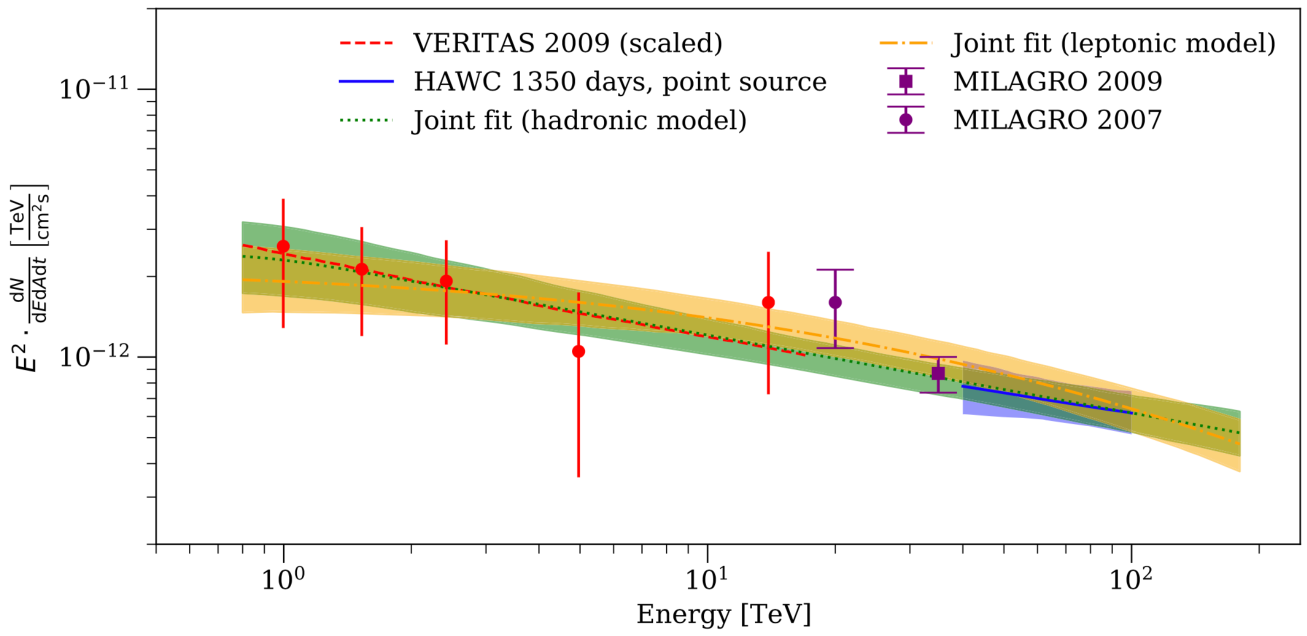
PeVatron Studies: The Boomerang Region



ApJL, Vol. 896, Is. 2, id.L29, 9 pp. (2020)

- Very-high-energy γ-ray emission > 100 TeV from HAWC J2227+610
- Excess well isolated and inconsistent with background fluctuations at the 6.2 σ level (pre-trials), or about 4.3 σ (post-trials considering HAWC's entire FoV)
- Right figure:
  - Best-fit position of HAWC J2227+610 is consistent with the VHE detections by VERITAS and Milagro, and with the position of PSR J2229+6114 (within uncertainties)
  - Heat map: Molecular column density
  - Pink contours: 1.4 GHz continuum brightness temperature from the Canadian Galactic Plane Survey

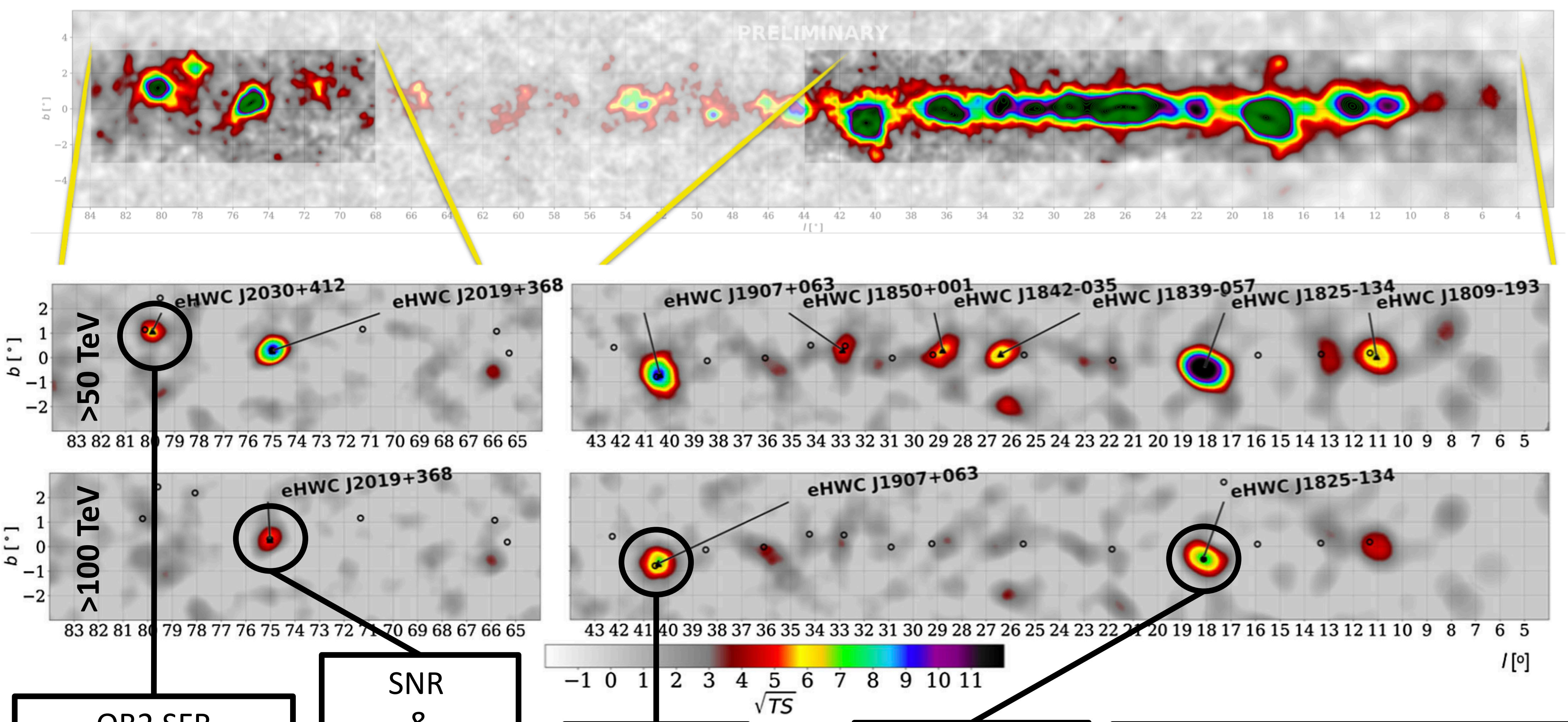
PeVatron Studies: The Boomerang Region



ApJL, Vol. 896, Is. 2, id.L29, 9 pp. (2020)

- Gaussian extent of HAWC J2227+610 is constrained to be $< \pm 0.232°$, morphology is consistent with VERITAS
- Joint VERITAS–HAWC spectrum well fit by a power law ($\gamma \approx -2.3$) from ~ 0.9 to ~ 180 TeV:
  - Emission can be interpreted to be originating from protons with a lower limit in their cutoff energy of 800 TeV.
  - Most likely source of the protons: the associated supernova remnant G106.3+2.7
  - But purely leptonic origin of the observed emission cannot be excluded at this time

PeVatron Searches: Recent More Detailed Results



PRL 124, 021102 (2020)

OB2 SFR (Cygnus Cocoon)*

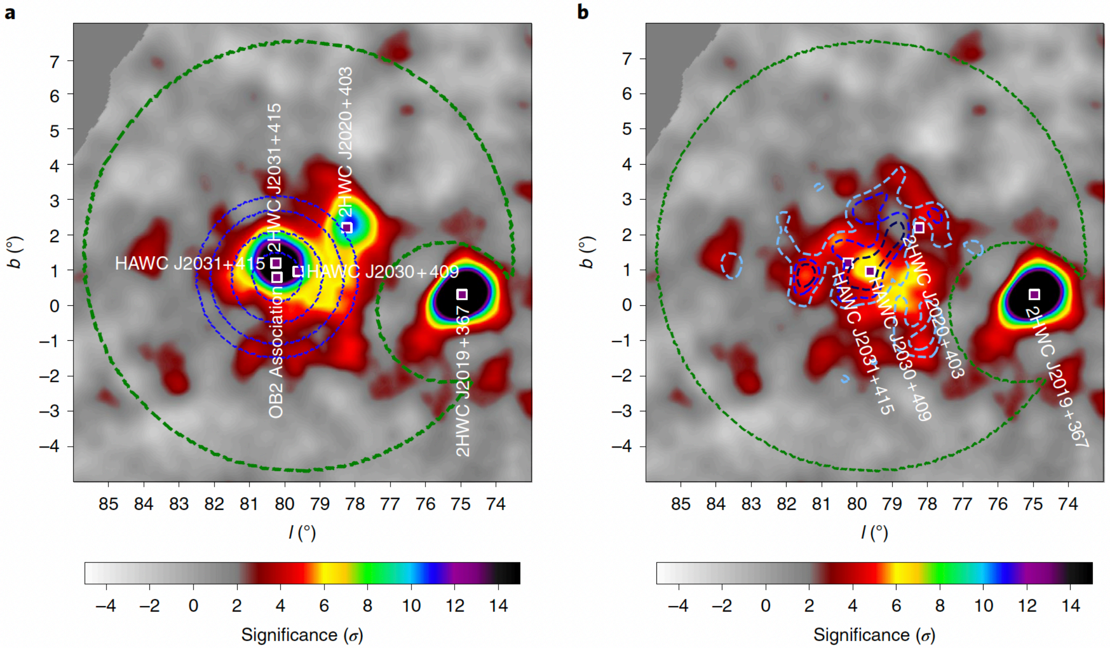
SNR & PWN (Dragon Fly)

PWN/Halo, SNR

PWN, UnID + GMC**

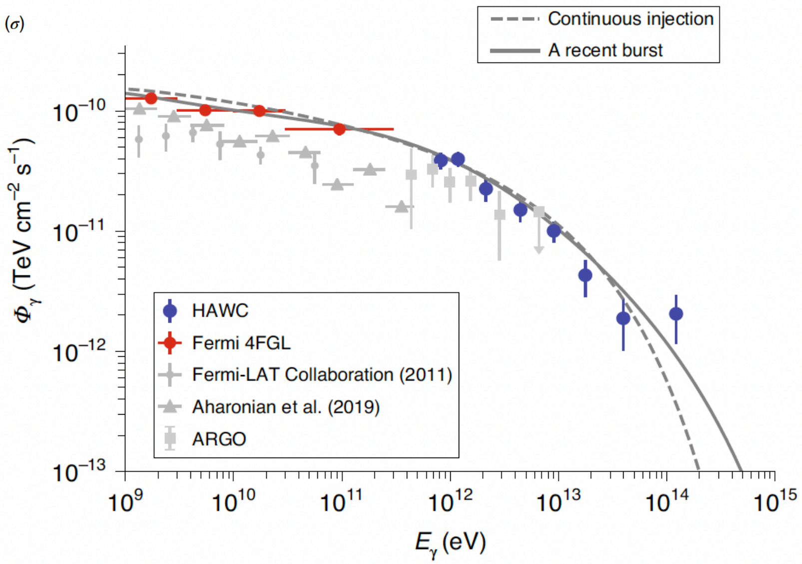
all these analyses involve multi-component fits
And have been published

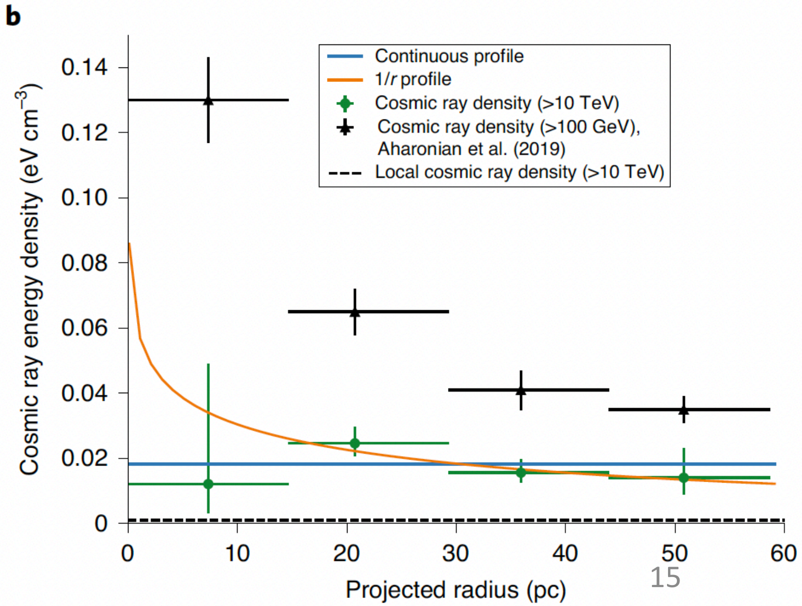
PeVatron Searches: An Example



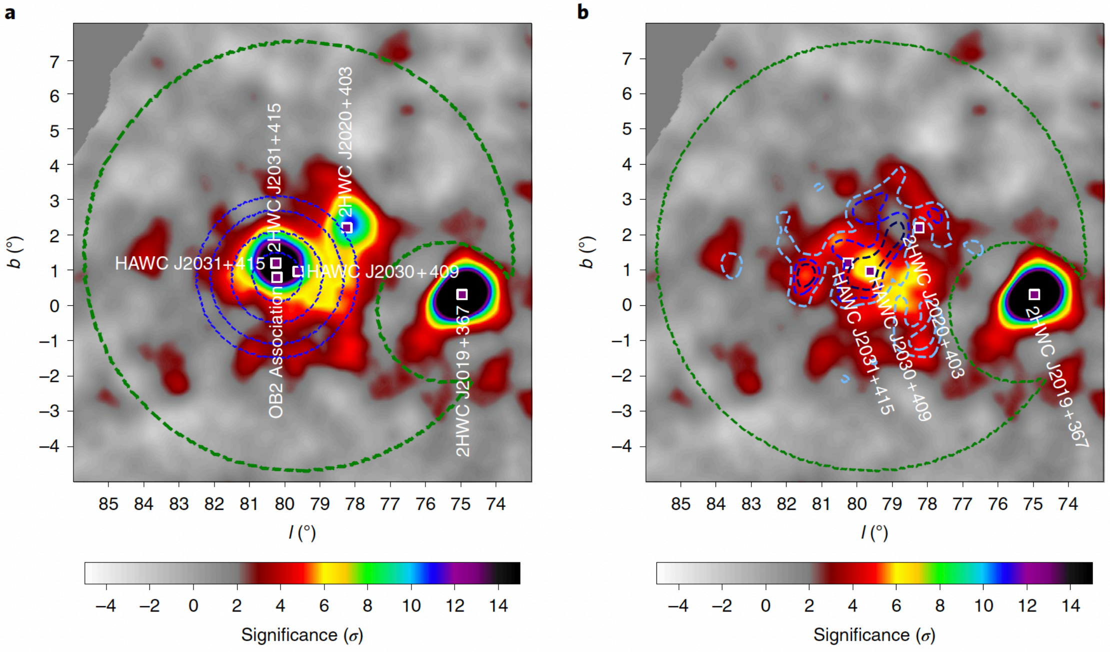
OB2 SFR
(Cygnus Cocoon)

[Nat. Astro., 5, 465-471 (2021)]





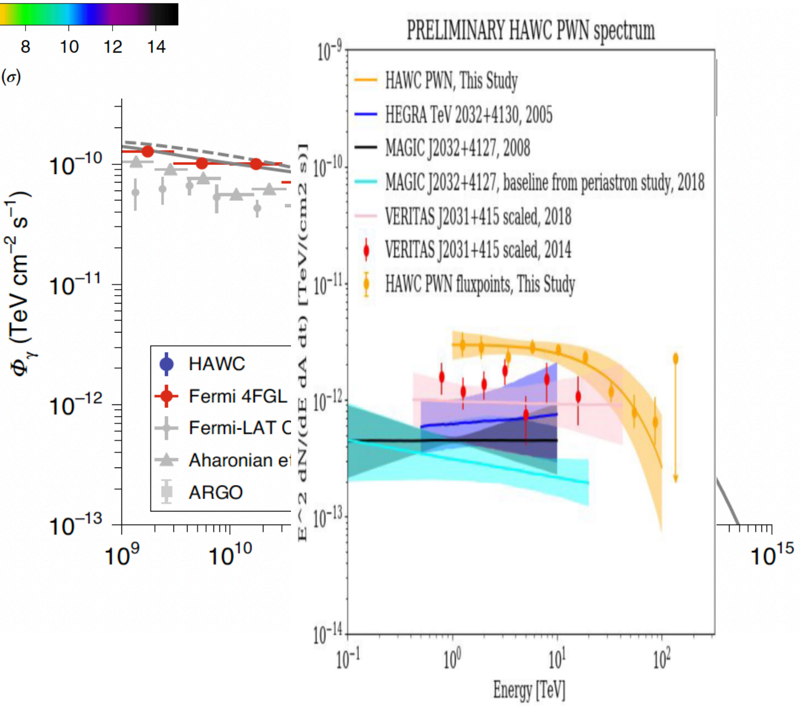
PeVatron Searches: An Example

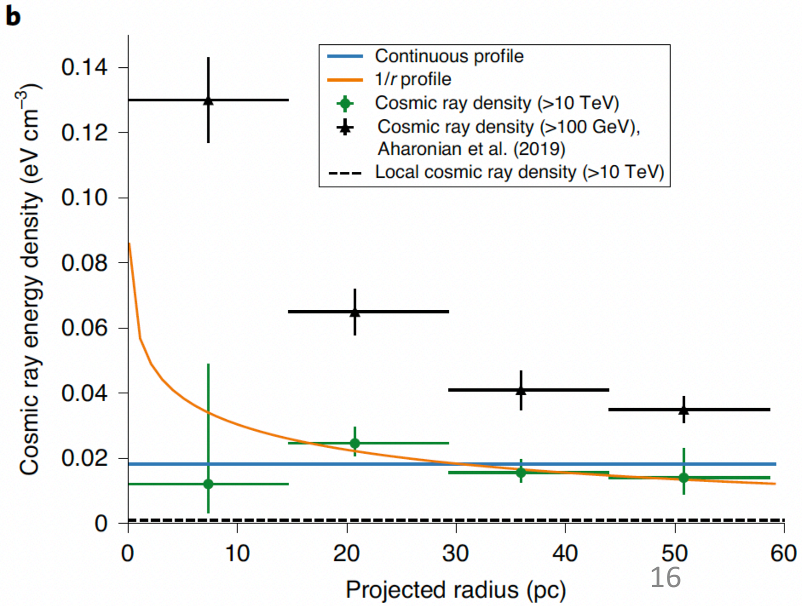


OB2 SFR
(Cygnus Cocoon)

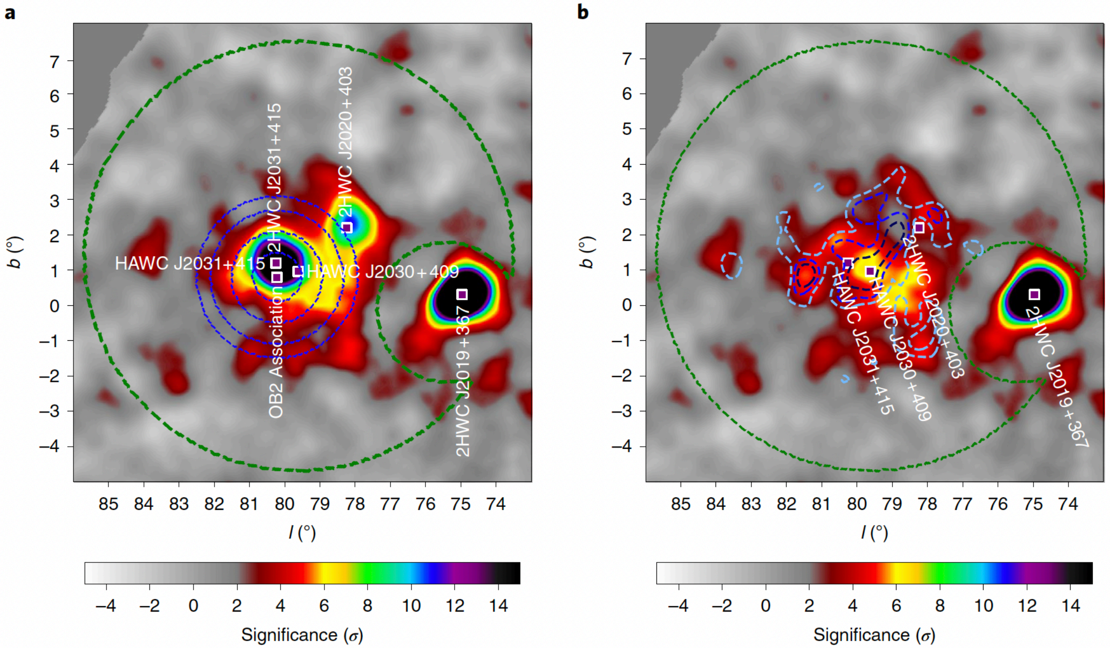
Nat. Astro., 5, 465-471 (2021)

PoS ICRC2021, 836 (2021)





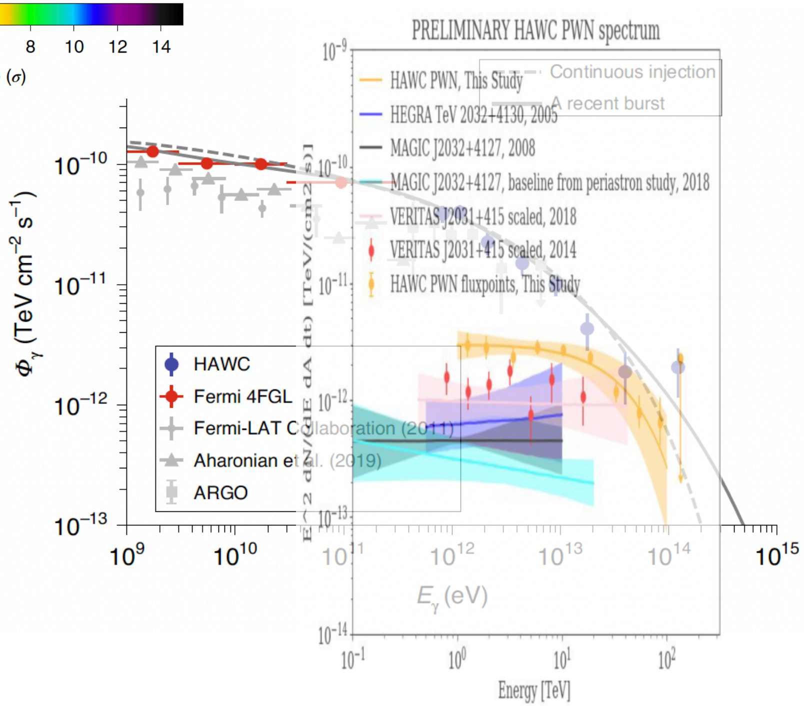
PeVatron Searches: And Example

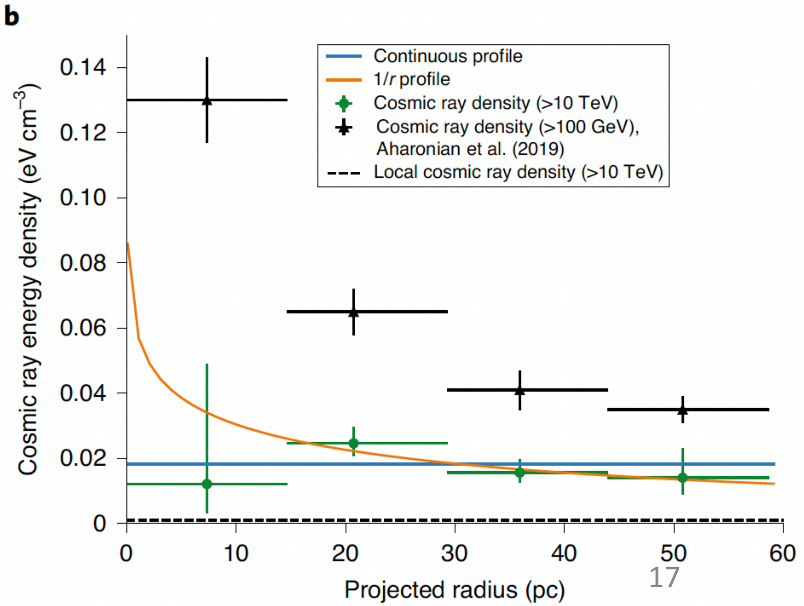


OB2 SFR
(Cygnus Cocoon)

Nat. Astro., 5, 465-471 (2021)

PoS ICRC2021, 836 (2021)



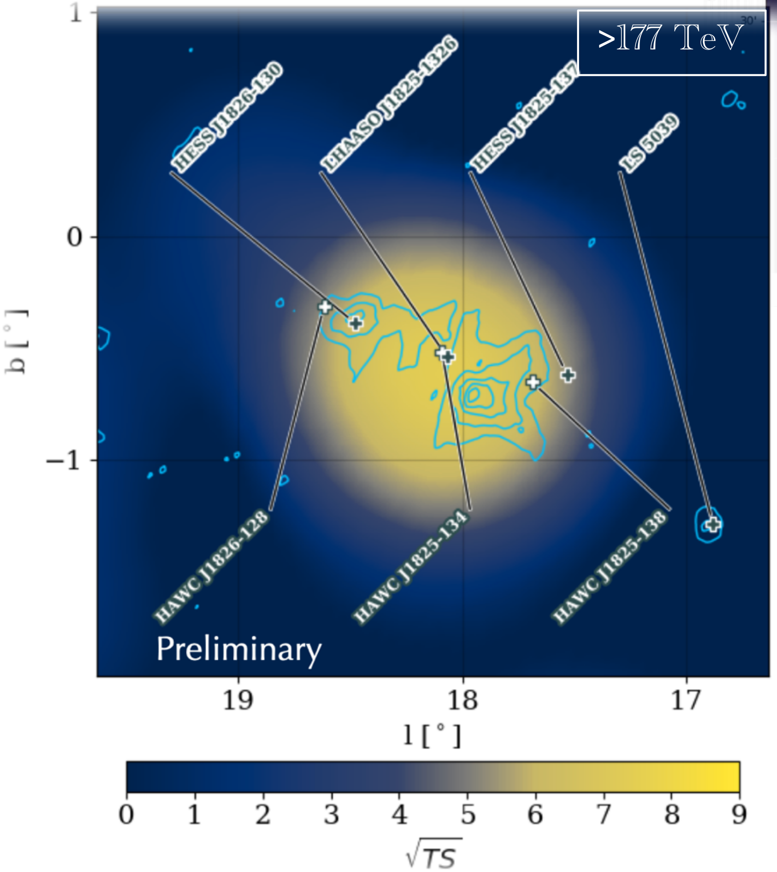


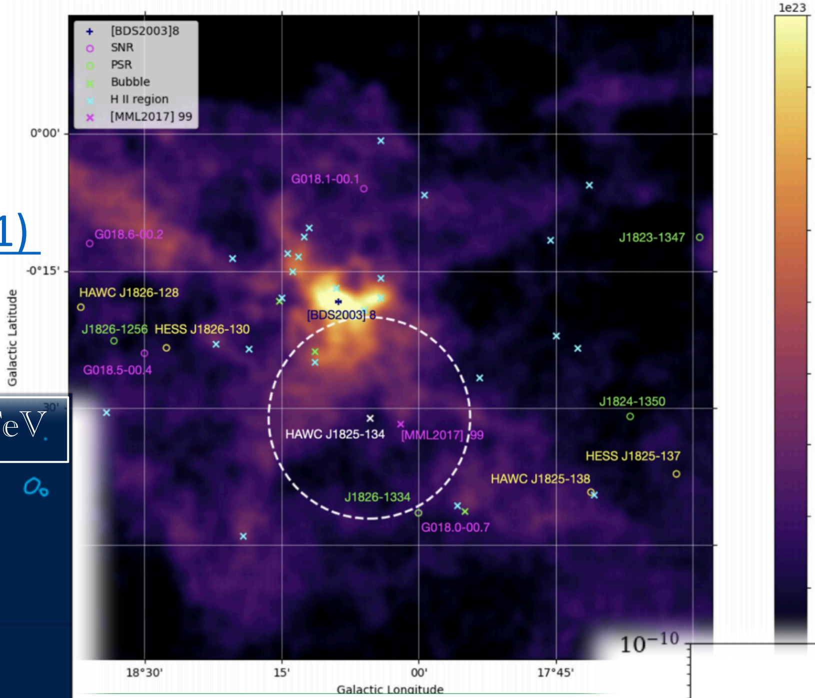
PeVatron Searches: Another Example

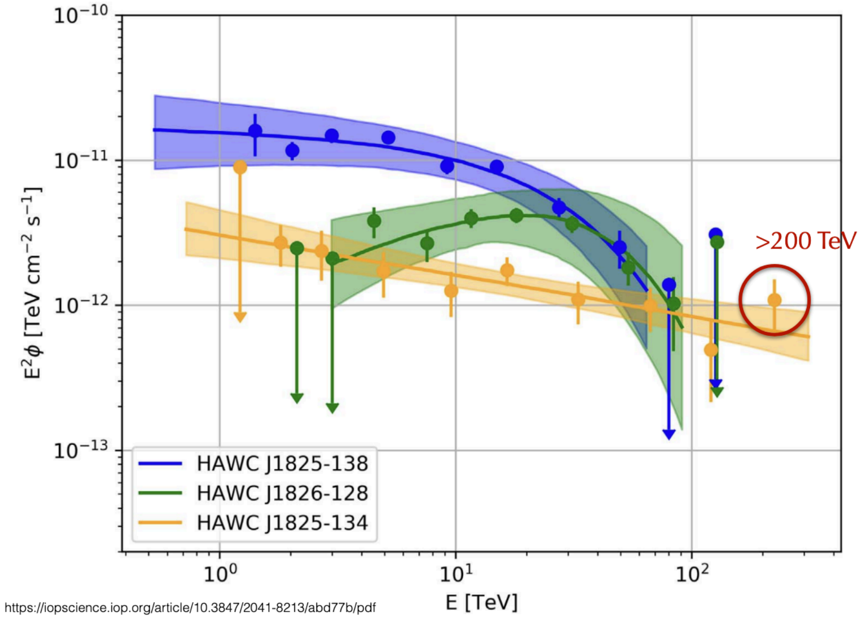
PWN,
UnID + GMC

[ApJL, Vol. 907, Is. 2, id. L30, 9 (2021)](https://iopscience.iop.org/article/10.3847/2041-8213/abd77b/pdf)

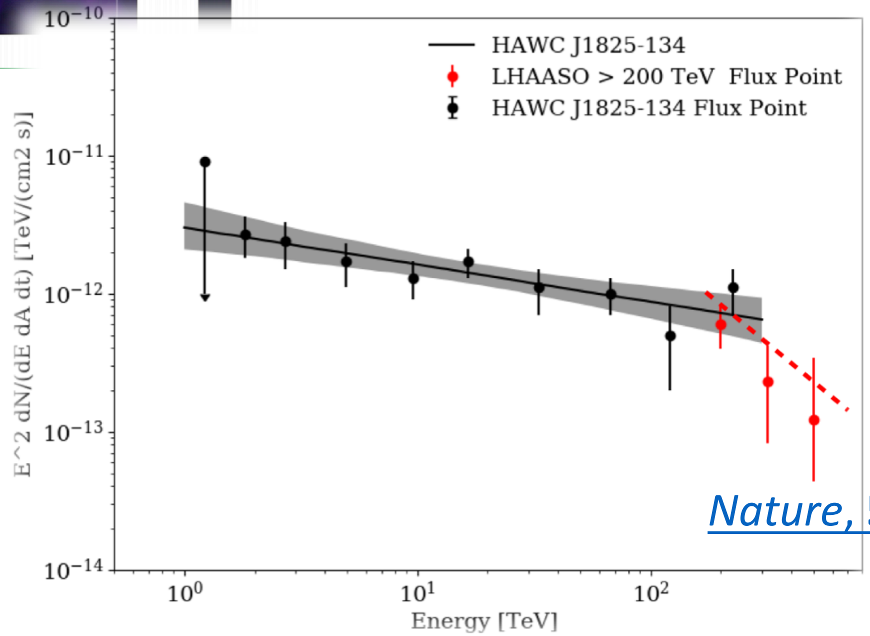
PoS ICRC2021, 840 (2021)







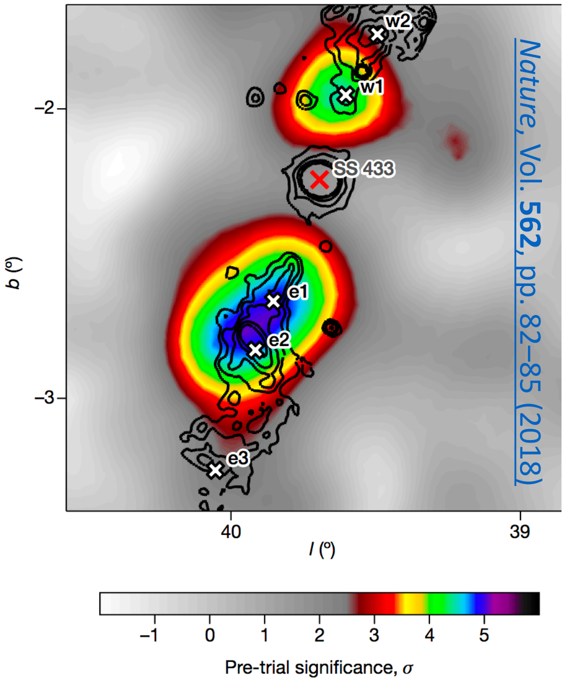
https://iopscience.iop.org/article/10.3847/2041-8213/abd77b/pdf

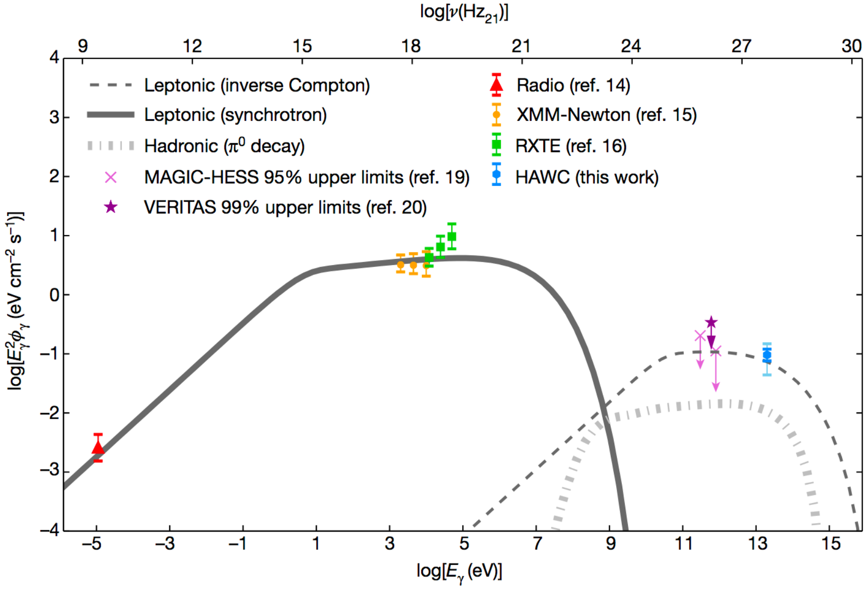


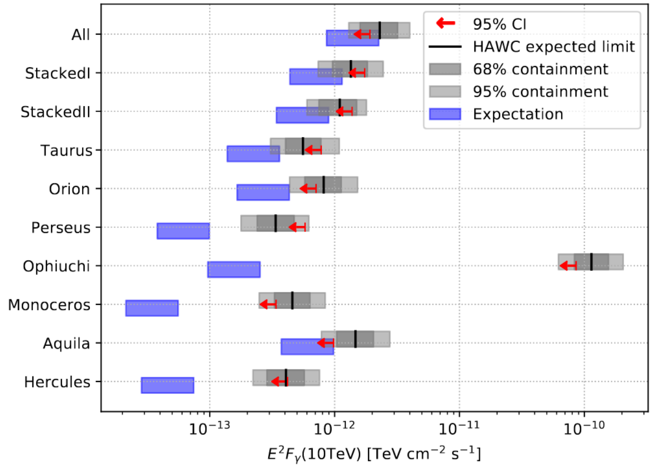
Nature, **594**, 33–36 (2021)

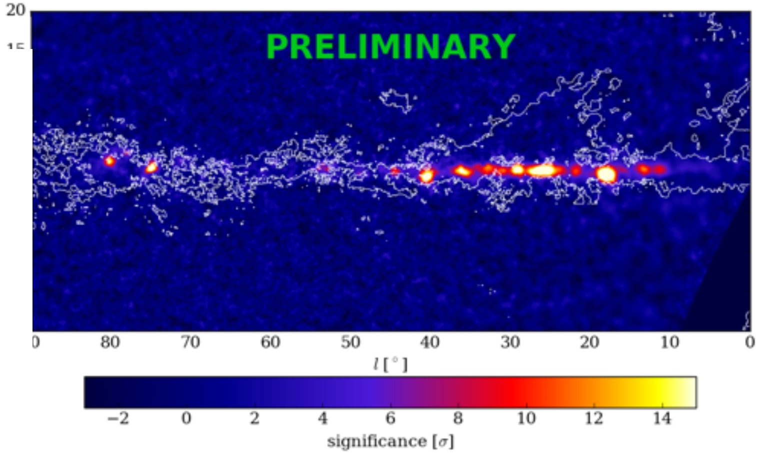
Other Recent Results of Probing our Galaxy

- ***Recent:***
  - 10s TeV $\gamma-$ray emission from ***SS433***; particle acceleration is occurring in the jets, not in the central binary
  - Upper Limits on ***Fermi Bubble*** emission disfavoring proton injection spectrum that extends >100 TeV w/o suppression [[ApJ, Vol. 842, Is. 2, id. 85, 9 pp. (2017)]]
  - Probe of the “sea” of cosmic rays in distant galactic regions through their interaction with ***Giant Molecular Clouds*** that generates multi-TeV $\gamma-$ray emission
- ***Forthcoming:***
  - Test of ***Galactic Diffuse Emission*** Models at multi-TeV (GALPROP, DRAGON)

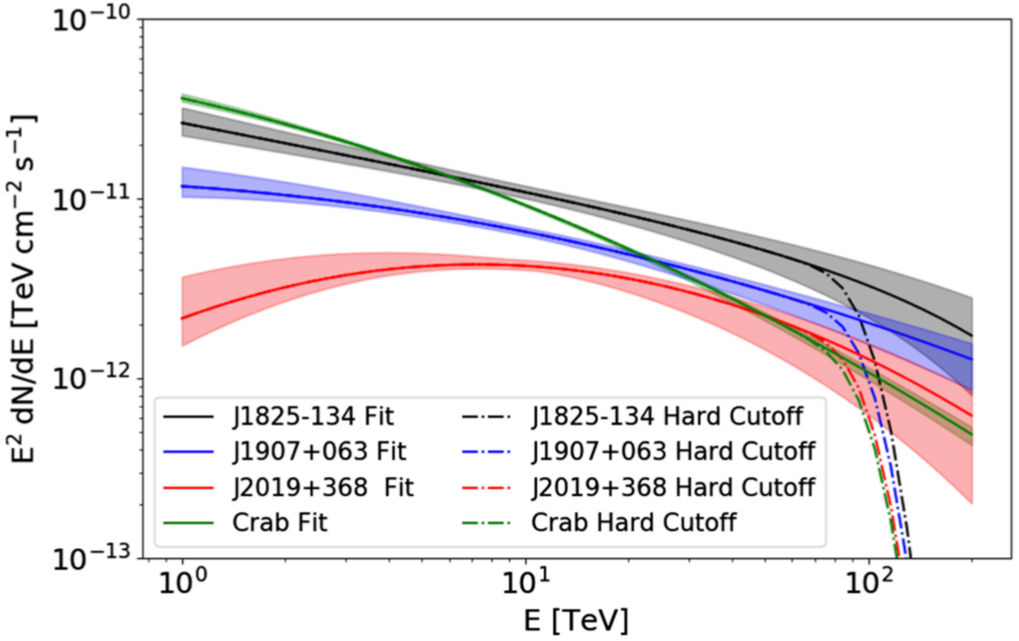


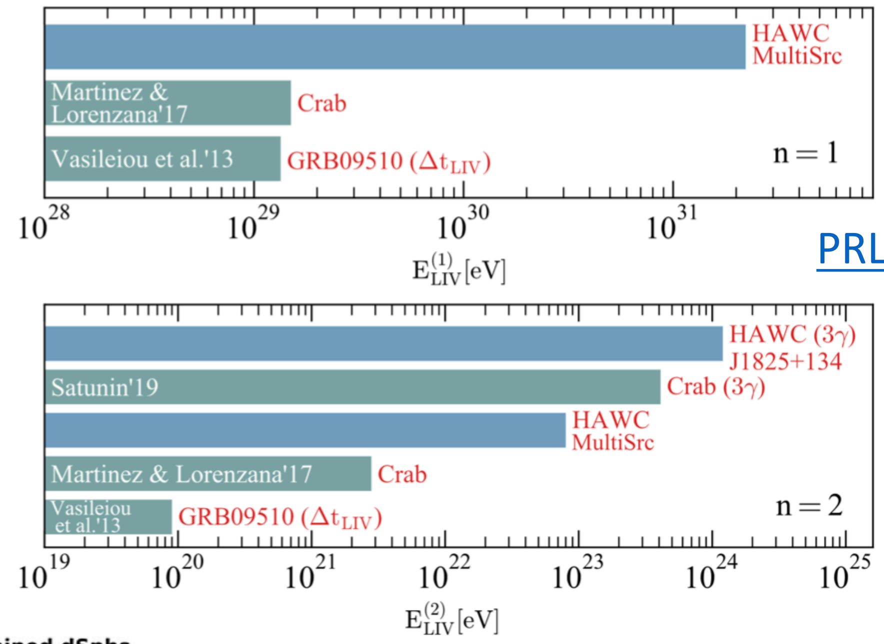




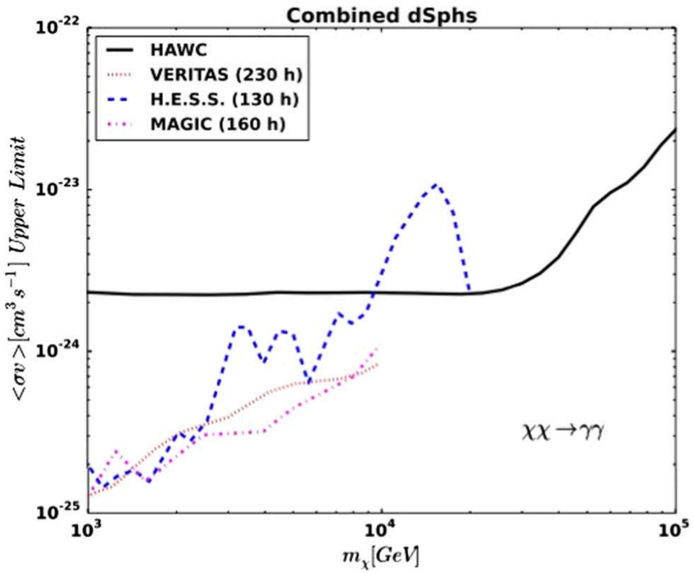


Fundamental Physics: LIV & Dark Matter





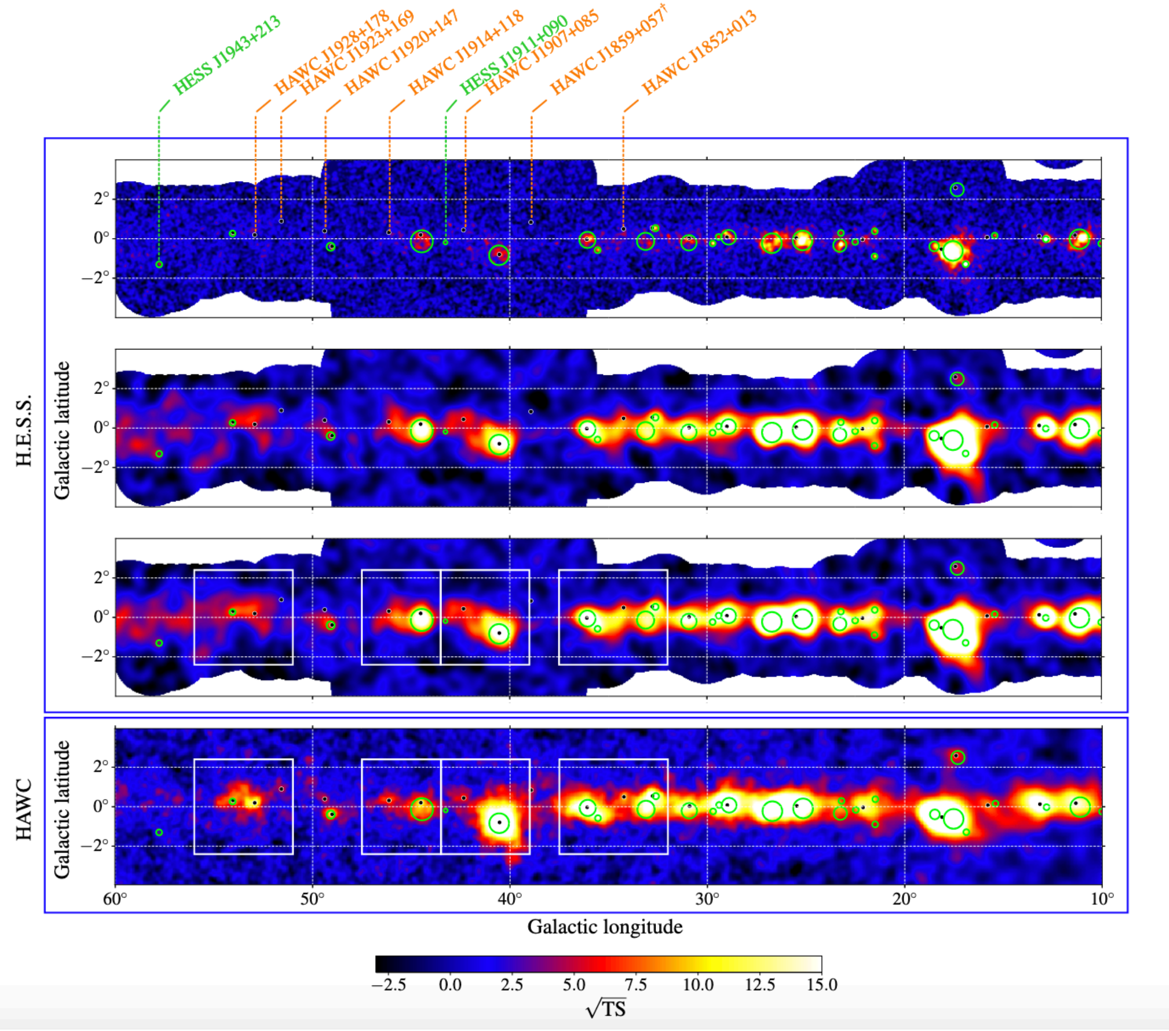
[PRL, 124, Is. 13, id.13110 (2020)](#)



[PRD, 101, Is. 10, article id.10300 (2020)](#)

Future Improvements

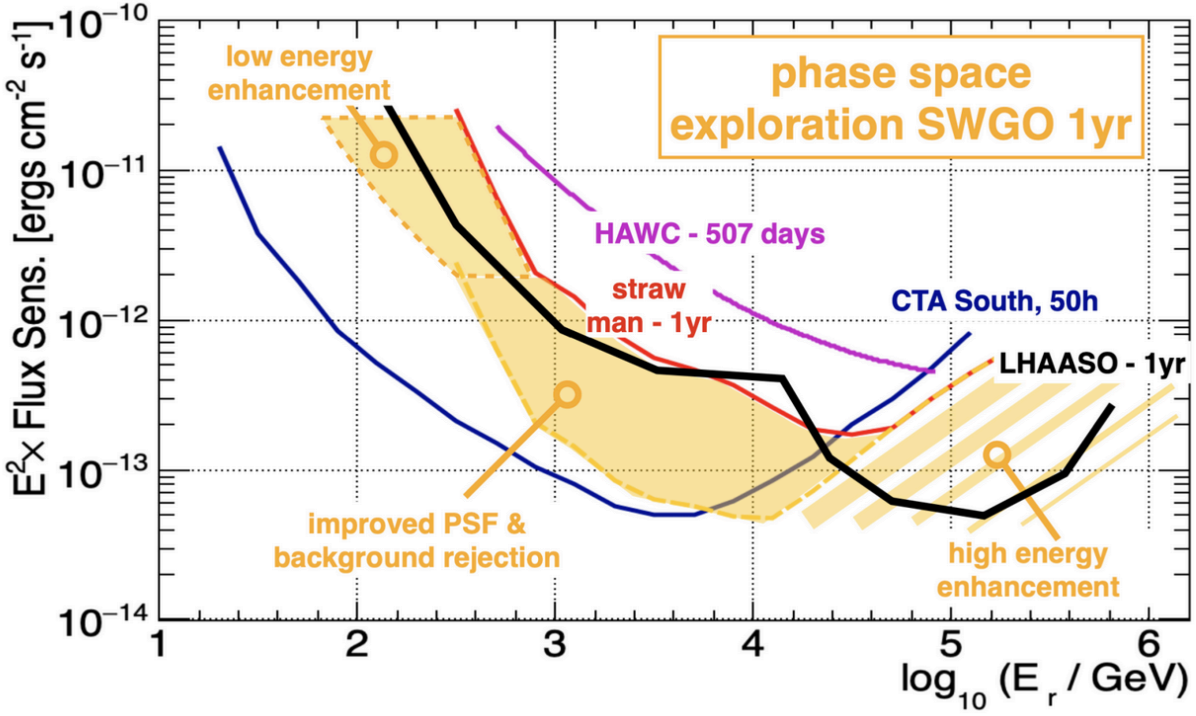
- Currently testing a new data pass of main array data (***pass5***)
  - Significant improvement of PSF (~0.1 deg at higest energy)
  - Significant improvements at low energies
  - Some forthcoming (unexpected?) results (over the next few months)
- Working on combining main array data with outrigger data (***pass6***)
- Continuing to work on combining HAWC data with data from other observatories (both electromagnetic and other messenger) and making HAWC data usable for the community



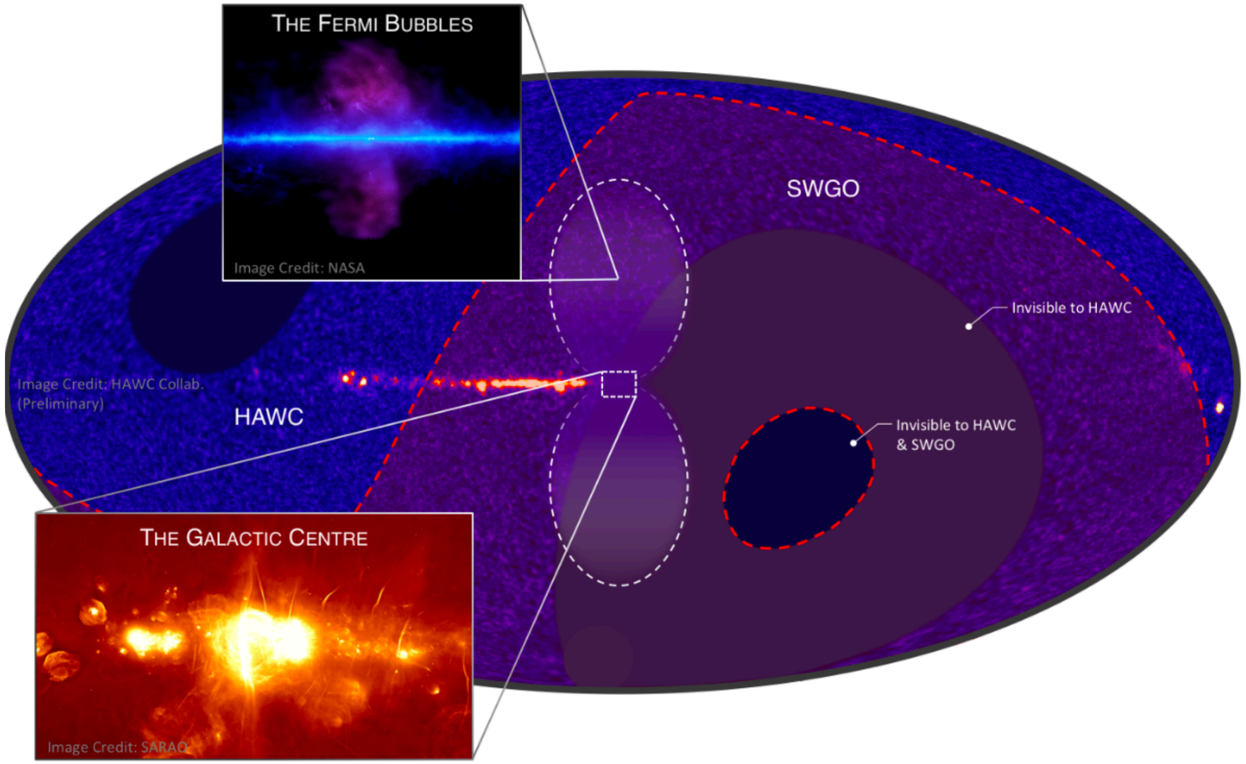
[ApJ, Vol.917, Is. 1, id.6, 16 pp. (2021)](ApJ, Vol.917, Is. 1, id.6, 16 pp. (2021))

The Future: Southern Wide Field Gamma-Ray Observatory*

SWGO will be located in Argentina, Bolivia, Chile, or Peru



[*PoS* ICRC2021, 903 (2021)]



[Astro2020: APC White Paper; BAAS, Vol. 51, Is. 7, 109 (2019)]

* (Endorsed by Astro2020 PAG committee)

Summary

- Ever more sensitive techniques used in **HAWC** data reconstruction & analysis
- With **HAWC and LHAASO,** we now have two instruments with unprecedented sensitivity > 10 TeV that also provide different complementary instantaneous views of the sky (time domain)
- Multi-Instrument and -messenger analyses will provide unprecedented science output