ACDM is alive and well!

A. Blanchard, J.-Y. Héloret, B.Lamine, S. Ilić, I.Tutusaus



Avignon, May 4th, 2023



\rightarrow predictive

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 $\rightarrow {\rm predictive}$

 \rightarrow accurate parameters determination \sim % precision.

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Tensions.

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Tensions. *H*₀

 $\rightarrow \text{ predictive}$

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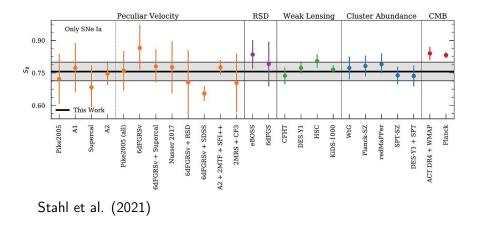
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Tensions.

 H_0

 S_8

The amplitude of matter fluctuations tension, i.e. S_8 tension.



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weak lensing

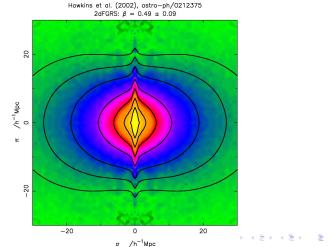
- weak lensing
- cluster abundance

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- cluster abundance + mass calibration (1 b)

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• RSD (redshift space distorsion)

- weak lensing
- cluster abundance + mass calibration (1 b)
- RSD (redshift space distorsion) $| \rightarrow f \sigma_8 |$



Next step: SZ Clusters & eBOSS RSD

Recipe:

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Recipe:

• use only "local" data i.e. $z \ll 1000$

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• work in the ACDM framework.

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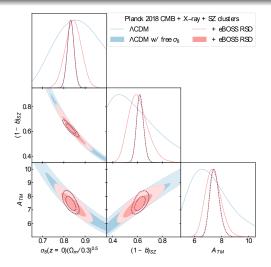
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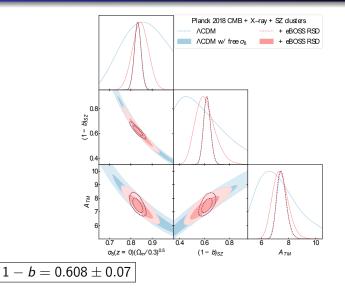
- use only "local" data i.e. $z \ll 1000$
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- Planck SZ clusters counts
- 1 b as a free parameter.
- Cosmological parameters from Planck CMB but σ_8

Constraining 1-b, σ_8 , S_8

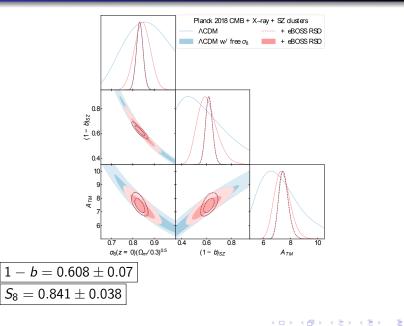


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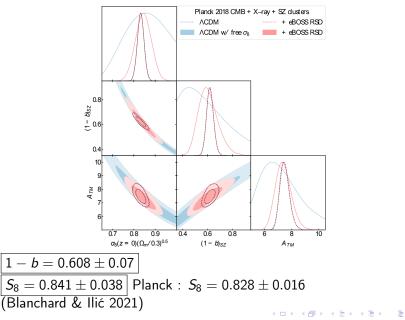
Constraining 1 - b, σ_8 , S_8



Constraining 1 - b, σ_8 , S_8



Constraining 1 - b, σ_8 , S_8



New cluster calibration: ACT DR5

arXiv:2304.10219

$1-b = 0.65 \pm 0.05$

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arXiv:2304.10219

$$1-b = 0.65 \pm 0.05$$

This translates to :

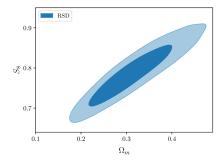
$$S_8 = 0.818 \pm 0.027$$

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One step further: RSD from surveys

| Survey | \mathbf{z} | $\mathrm{f}\sigma_8$ | Refs |
|--------------|--------------|----------------------|------|
| 2MFT | 0.001 | 0.51 + / -0.085 | [19] |
| 6dFGS | 0.067 | 0.423 + / -0.055 | [20] |
| SDSS DR13 | 0.1 | 0.48 + / -0.16 | [21] |
| 2dFGRS | 0.17 | 0.51 + / -0.06 | [22] |
| GAMA | 0.18 | 0.36 + / - 0.09 | [23] |
| WiggleZ | 0.22 | 0.42 + / -0.07 | [24] |
| SDSS LRG60 | 0.25 | 0.35 + / - 0.06 | [25] |
| BOSS LOW Z | 0.32 | 0.48 + / - 0.1 | [26] |
| GAMA | 0.36 | 0.44 + / - 0.06 | [23] |
| SDSS LRG 200 | 0.37 | 0.46 + / - 0.04 | [25] |
| WiggleZ | 0.41 | 0.45 + / -0.04 | [24] |
| CMASS BOSS | 0.57 | 0.453 + / -0.02 | [27] |
| WiggleZ | 0.6 | 0.43 + / -0.04 | [24] |
| VIPERS | 0.6 | 0.48 + / -0.12 | [28] |
| SDSS IV | 0.69 | 0.447 + / -0.039 | [29] |
| VIPERS | 0.76 | 0.44 + / -0.04 | [30] |
| SDSS IV | 0.77 | 0.432 + / -0.038 | [31] |
| WiggleZ | 0.78 | 0.38 + / -0.04 | [24] |
| SDSS IV | 0.85 | 0.52 + / -0.10 | [32] |
| VIPERS | 0.86 | 0.48 + / -0.10 | [28] |
| SDSS IV | 0.978 | 0.379 + / -0.176 | [31] |
| SDSS IV | 1.23 | 0.385 + / - 0.1 | [31] |
| Fastsound | 1.4 | 0.494 + / -0.123 | [33] |
| SDSS IV | 1.52 | 0.426 + /-0.077 | [34] |
| SDSS IV | 1.944 | 0.364 + / -0.106 | [31] |

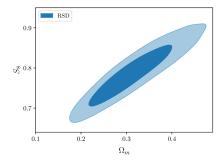
RSD from surveys: constraints



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RSD from surveys: constraints

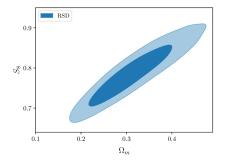


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Not surprisingly strong degeneracy

RSD from surveys: constraints

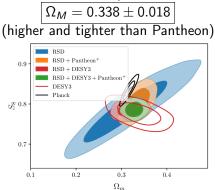


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Not surprisingly strong degeneracy Need to combine with other low - z data

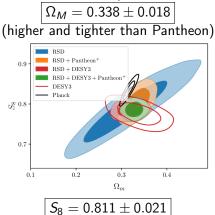
Pantheon+: SNIa Hubble diagram (Brout et al., 2022), for Λ CDM): $\Omega_M = 0.338 \pm 0.018$

Pantheon+: SNIa Hubble diagram (Brout et al., 2022), for ACDM):



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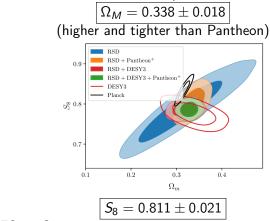
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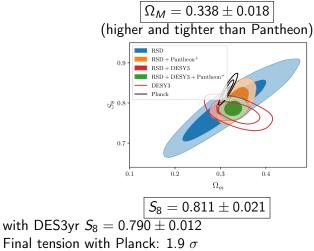
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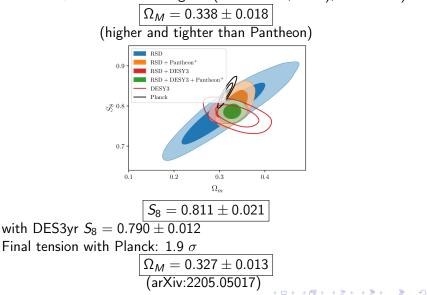
with DES3yr $S_8 = 0.790 \pm 0.012$

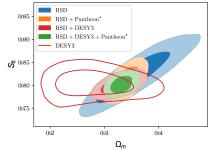
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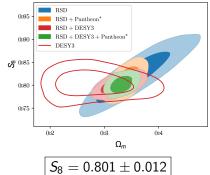
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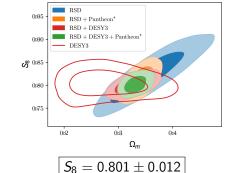
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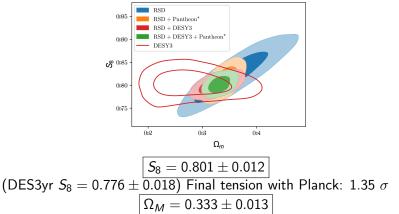


(DES3yr $S_8 = 0.776 \pm 0.018$) Final tension with Planck: 1.35 σ

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Last word on S_8 : ACT DR5

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DR6 Gravitational Lensing Map and Cosmological Parameters (arXiv:2304.05203v1)

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DR6 Gravitational Lensing Map and Cosmological Parameters (arXiv:2304.05203v1)

 $S_8 = 0.831 \pm 0.023$

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DR6 Gravitational Lensing Map and Cosmological Parameters (arXiv:2304.05203v1)

$$S_8 = 0.831 \pm 0.023$$

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Conclusion:

DR6 Gravitational Lensing Map and Cosmological Parameters (arXiv:2304.05203v1)

$$S_8 = 0.831 \pm 0.023$$

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Conclusion:

Four pieces of evidence for low z $S_8 \sim 0.81 \ (\pm 0.01)$

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DR6 Gravitational Lensing Map and Cosmological Parameters (arXiv:2304.05203v1)

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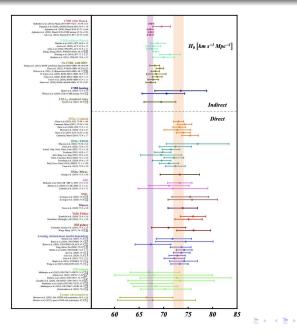
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Conclusion:

Four pieces of evidence for low z $S_8 \sim 0.81~(\pm 0.01)$

Consistent with LCDM Planck normalized!

H_0 tension



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Measuring the Tension

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$$\chi^2 = \sum \frac{(H_0 - \alpha_i \times H_{0,i})^2}{\sigma_i^2} \tag{1}$$

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With H_0 from **SH0ES**, **TF**, **SBF**, CCHP, MCP, Miras, BAO, Planck

$$\chi^2 = \sum \frac{(H_0 - \alpha_i \times H_{0,i})^2}{\sigma_i^2} \tag{1}$$

With H_0 from **SH0ES**, **TF**, **SBF**, CCHP, MCP, Miras, BAO, Planck

Akaike Information Criterium (AIC):

$$\Delta \text{AIC} = \Delta \chi^2 + 2\Delta p \,. \tag{2}$$

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for model comparison.

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| Model | χ^2 | ΔΑΙΟ | |
|-------|----------|------|--|
| ΛCDM | 37.0 | - | |

| Model | χ^2 | ΔΑΙϹ | |
|---------|----------|-------|-------|
| ΛCDM | 37.0 | _ | |
| ACDM E1 | 17.3 | -17.7 | Riess |

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|---------|----------|-------|----------|
| ΛCDM | 37.0 | _ | |
| ACDM E1 | 17.3 | -17.7 | Riess |
| ACDM E2 | 6.7 | -26.3 | Cepheids |

| Model | χ^2 | ΔΑΙϹ | |
|---------|----------|-------|----------|
| ΛCDM | 37.0 | _ | |
| ACDM E1 | 17.3 | -17.7 | Riess |
| ACDM E2 | 6.7 | -26.3 | Cepheids |
| ACDM E3 | 34.4 | -0.6 | BAO |

| Model | χ^2 | ΔΑΙΟ | |
|---------|----------|--------|----------|
| ΛCDM | 37.0 | - | |
| ACDM E1 | 17.3 | -17.7 | Riess |
| ACDM E2 | 6.7 | -26.3 | Cepheids |
| ACDM E3 | 34.4 | -0.6 | BAO |
| ACDM E4 | 19.2 | -15.76 | Planck |

| Model | χ^2 | ΔΑΙϹ | |
|---------|----------|--------|-------------|
| ΛCDM | 37.0 | _ | |
| ACDM E1 | 17.3 | -17.7 | Riess |
| ACDM E2 | 6.7 | -26.3 | Cepheids |
| ACDM E3 | 34.4 | -0.6 | BAO |
| ACDM E4 | 19.2 | -15.76 | Planck |
| ACDM E5 | 4.3 | -30.7 | Planck +BAO |

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Conclusion: The E2 model is performing better than any alternative published model to solve the H_0 tension!

$\Omega_{\textit{M}}=0.327\pm0.013$

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$\Omega_{\textit{M}}=0.327\pm0.013$

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using SH0ES: $H_0 = 73.3 \pm 1.04 \text{ km/s/Mpc}$

$$\Omega_M = 0.327 \pm 0.013$$

using SH0ES: $H_0 = 73.3 \pm 1.04$ km/s/Mpc we can infer :

 $\omega_{\textit{M}}=0.1753\pm0.0069$

$$\Omega_M = 0.327 \pm 0.013$$

using SH0ES: $H_0 = 73.3 \pm 1.04$ km/s/Mpc we can infer :

 $\omega_{\textit{M}}=0.1753\pm0.0069$

compared to Planck (+ext):

$$\omega_M = 0.1425 \pm 0.0012$$

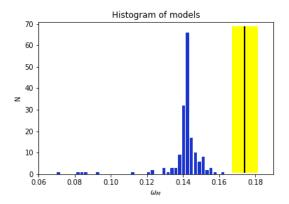
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4.7 σ away for ΛCDM

Let's take the \sim 200 models summarized in Di Valentino et al. (2021) In the realm of the Hubble tension – a review of solutions

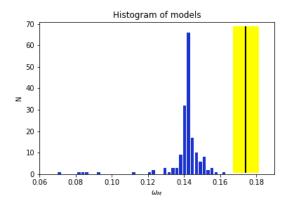
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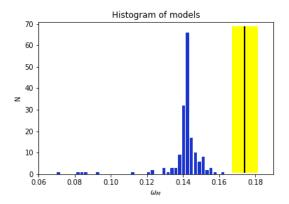
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Take the recent EDS model :

Let's take the \sim 200 models summarized in Di Valentino et al. (2021) In the realm of the Hubble tension – a review of solutions

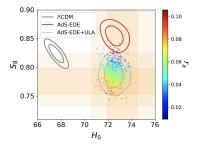


Take the recent EDS model : $\omega_M = 0.128 \pm 0.0037 \ 3\sigma$ away with this test (and all published EDE models).

arXiv:2107.13391v2: EDE+ ultralight axion

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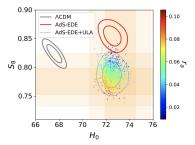
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However ω_m value > 2.6 σ away.

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Thank You ADDARD ED SOUCH