
Probing Dark Sectors with Evaporating Black Holes

Michael J. Baker

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28 September 2022

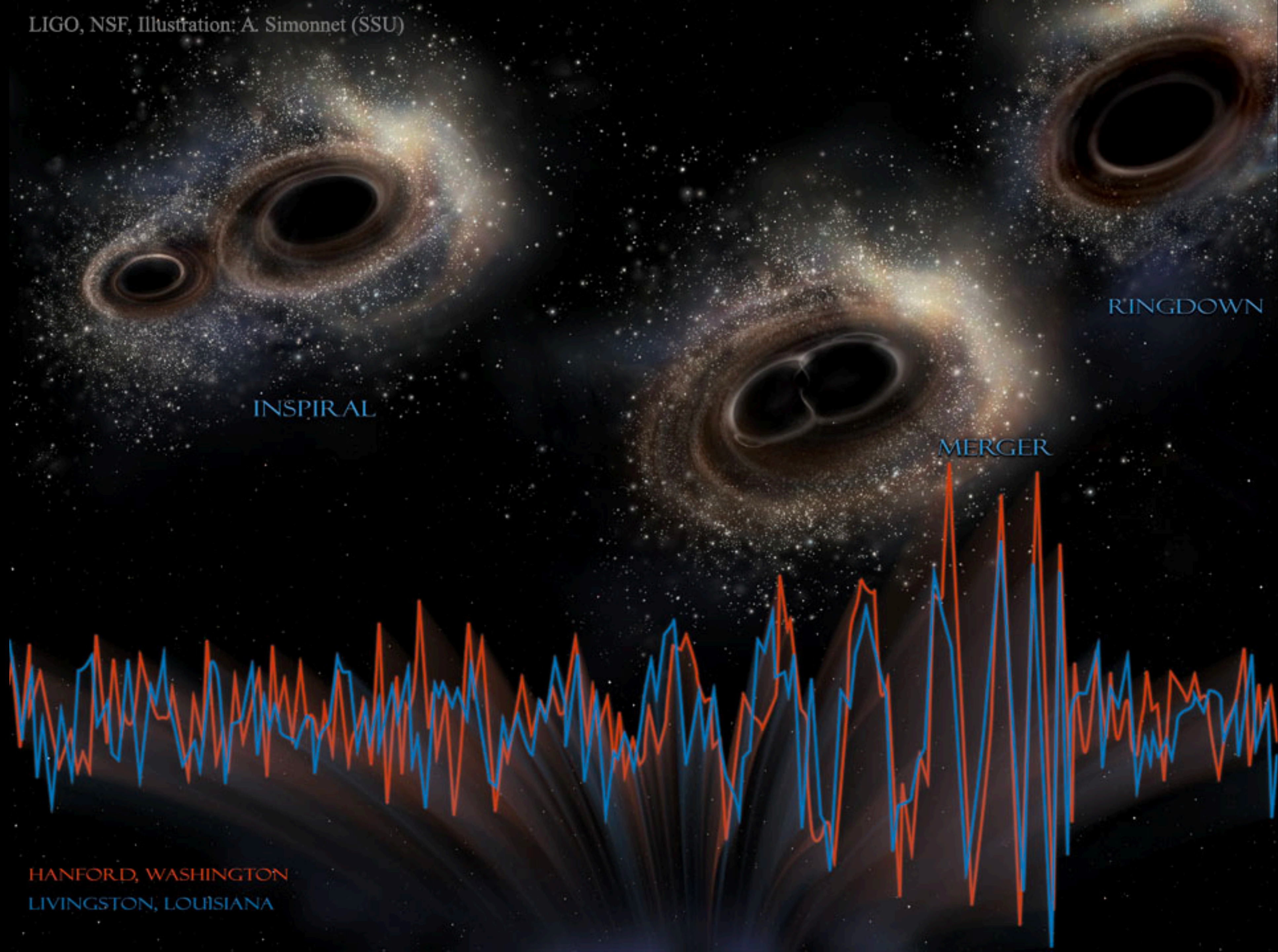
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Introduction



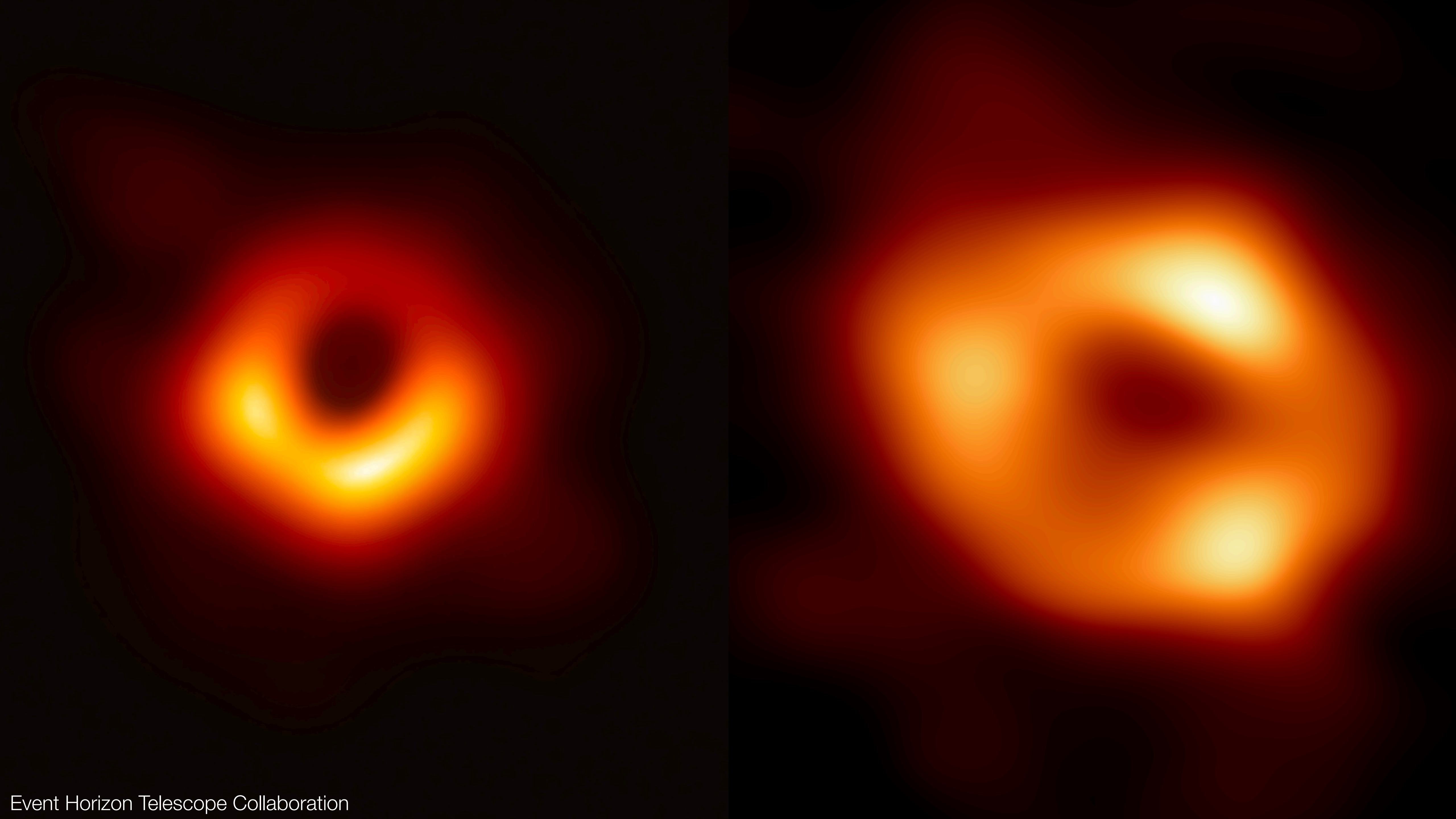
INSPIRAL

RINGDOWN

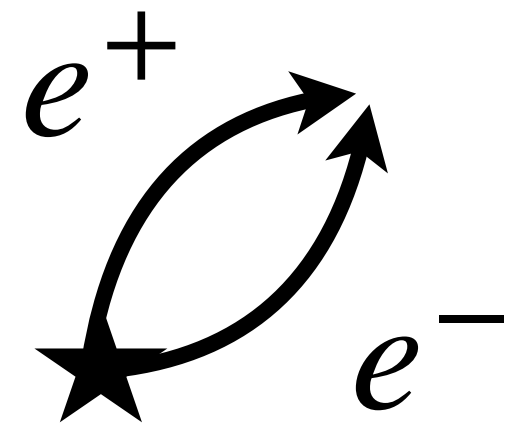
MERGER

HANFORD, WASHINGTON
LIVINGSTON, LOUISIANA

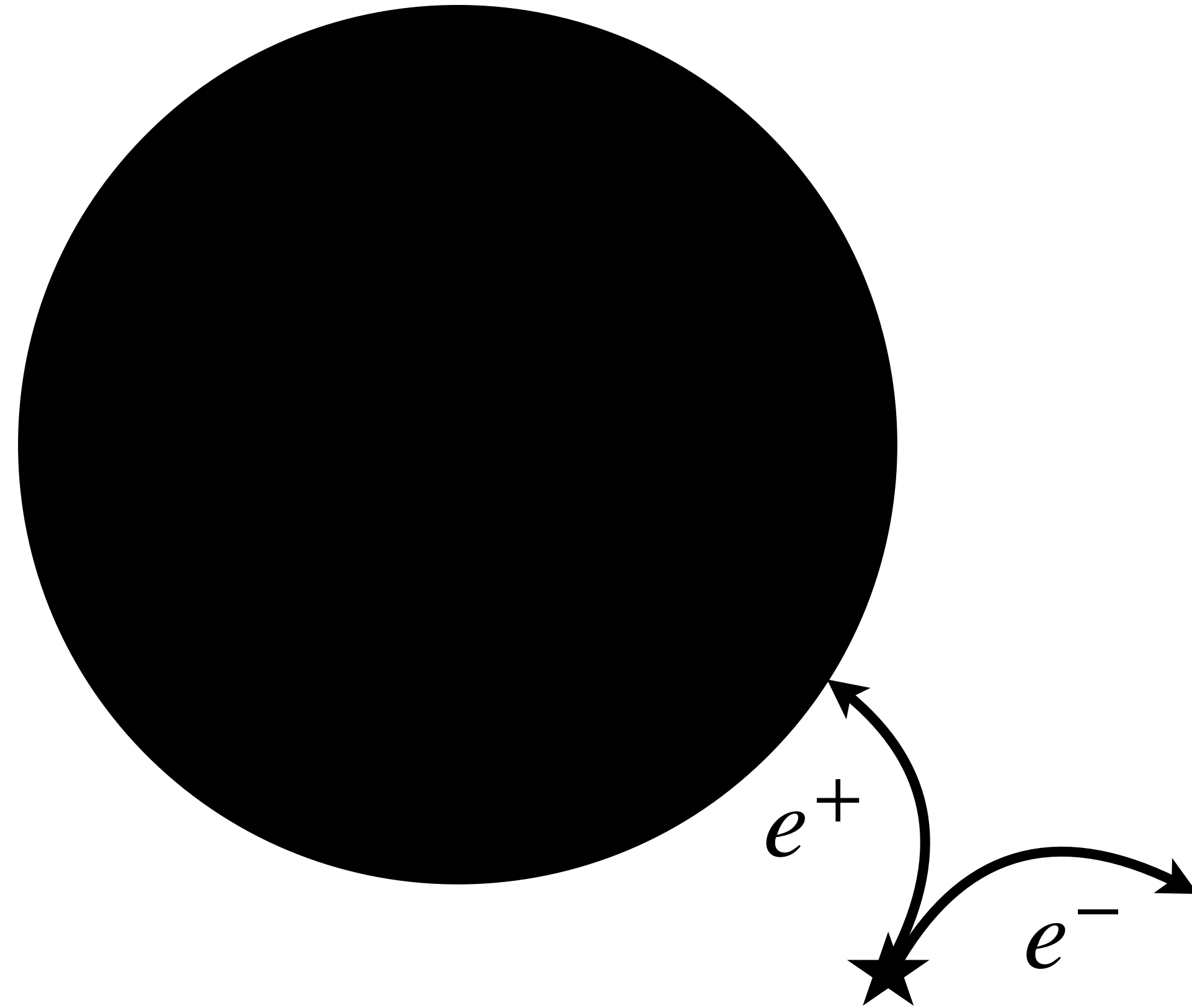
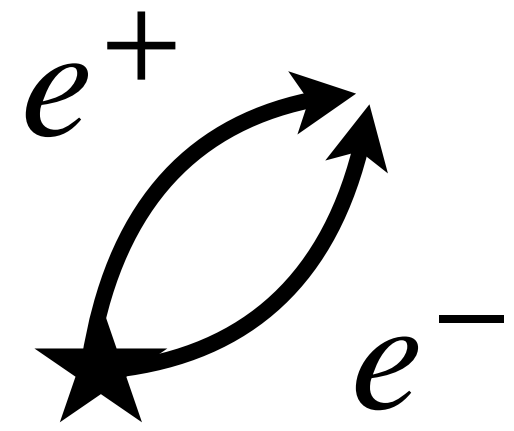




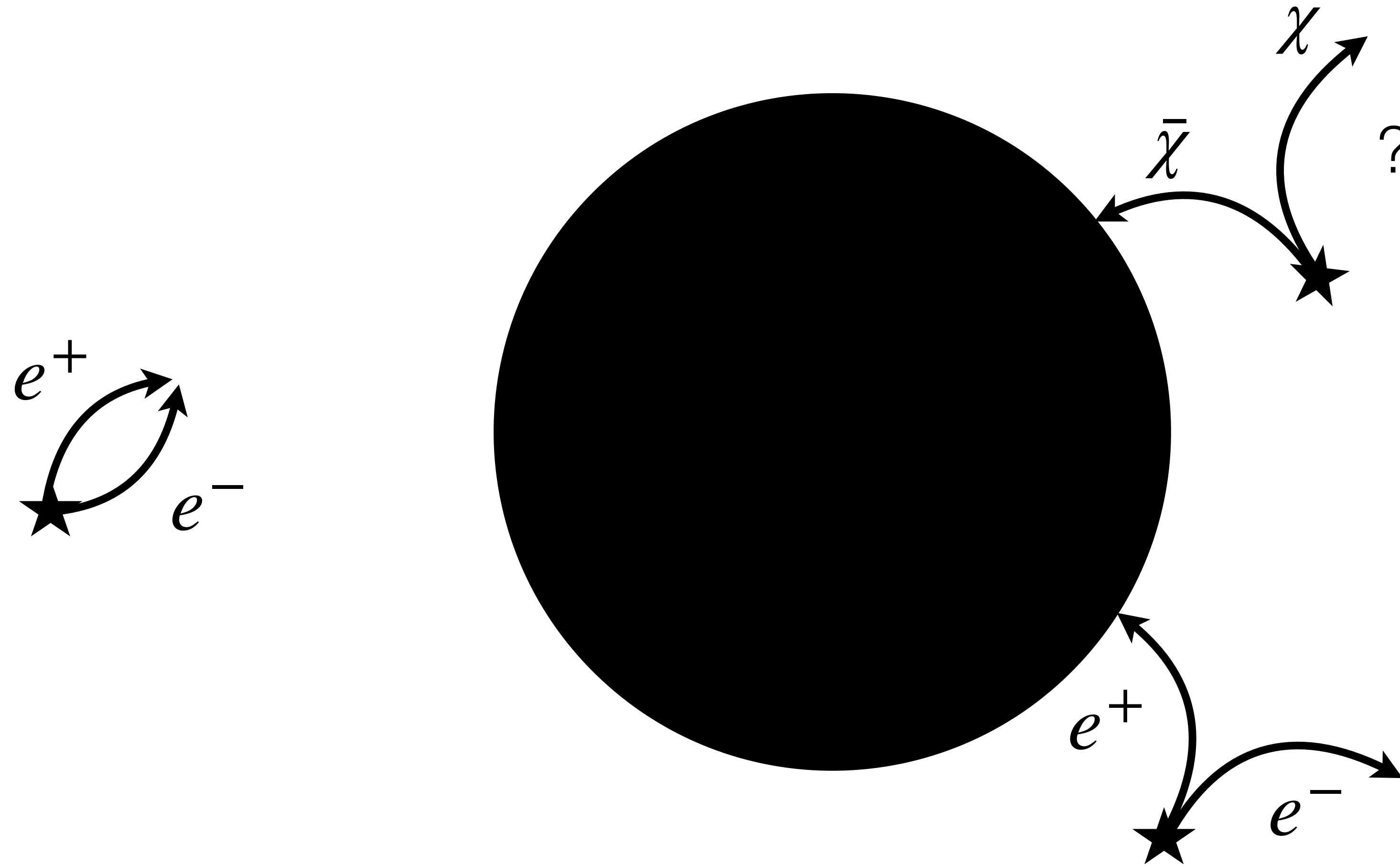
Hawking Radiation



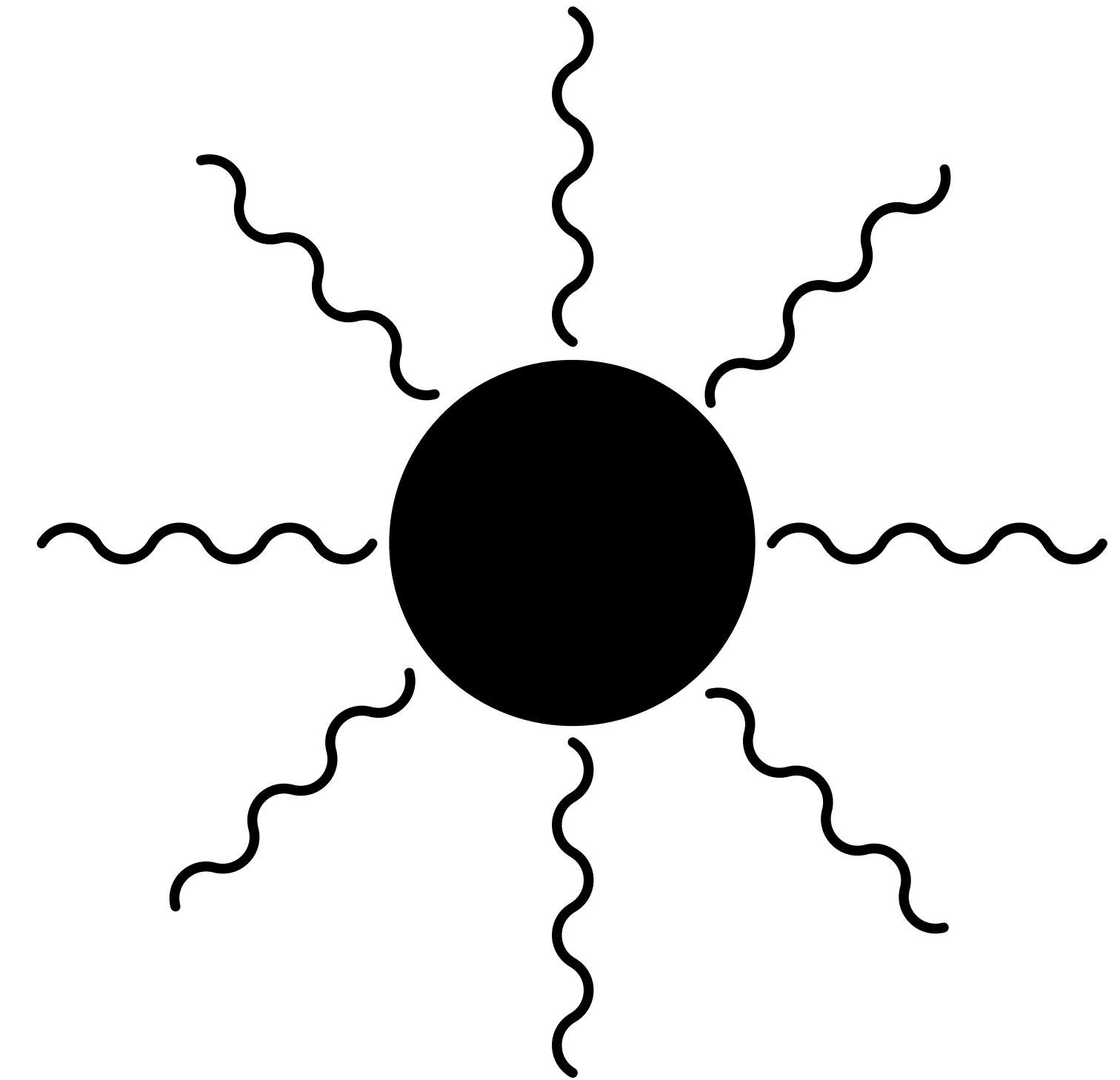
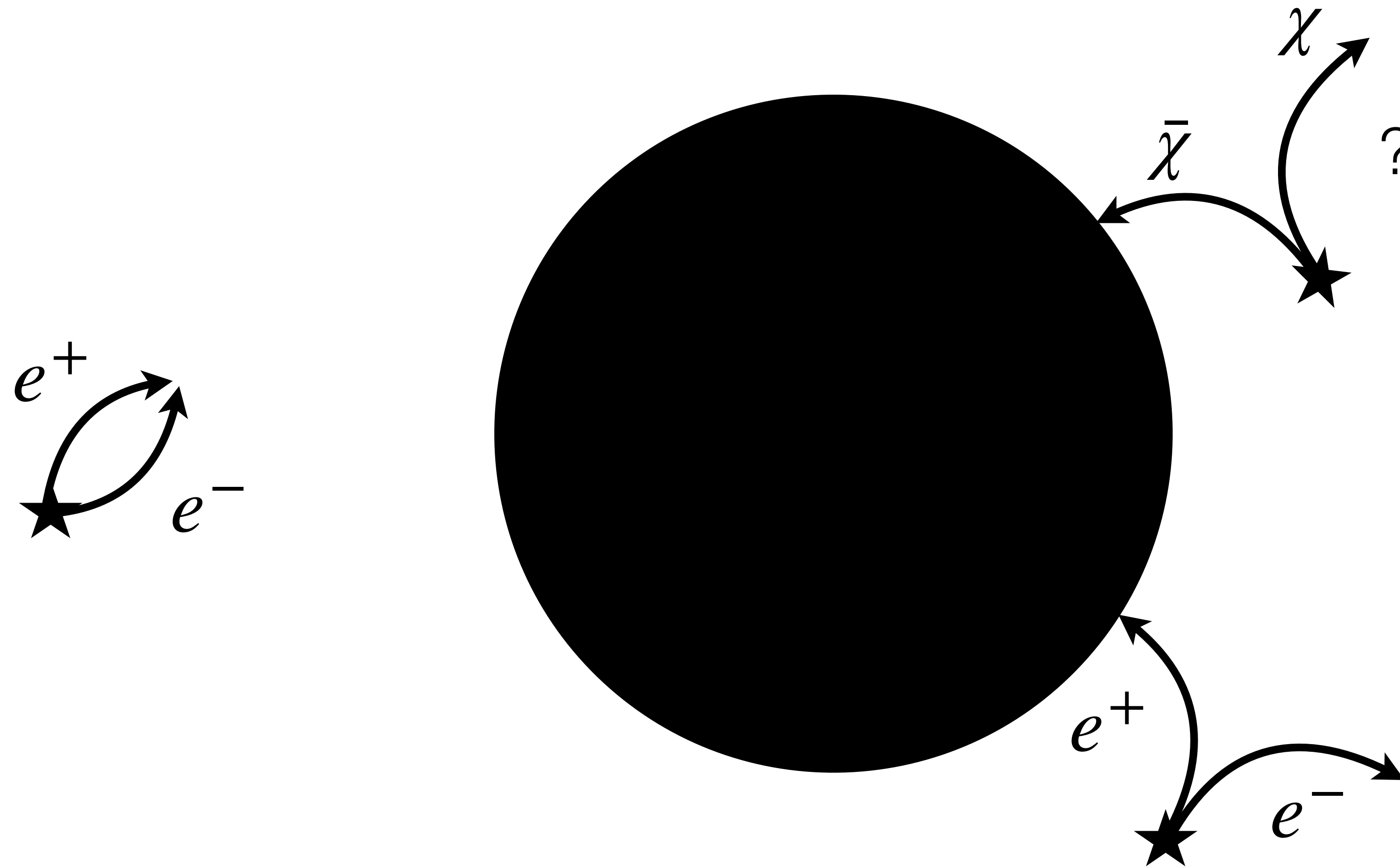
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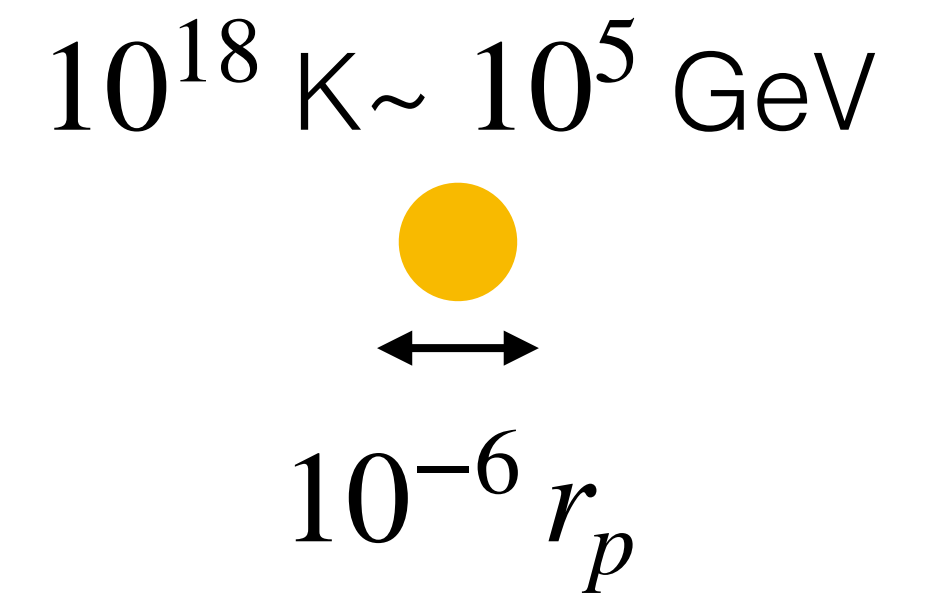
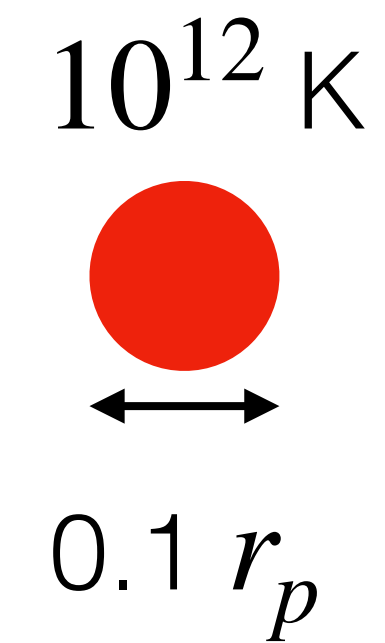
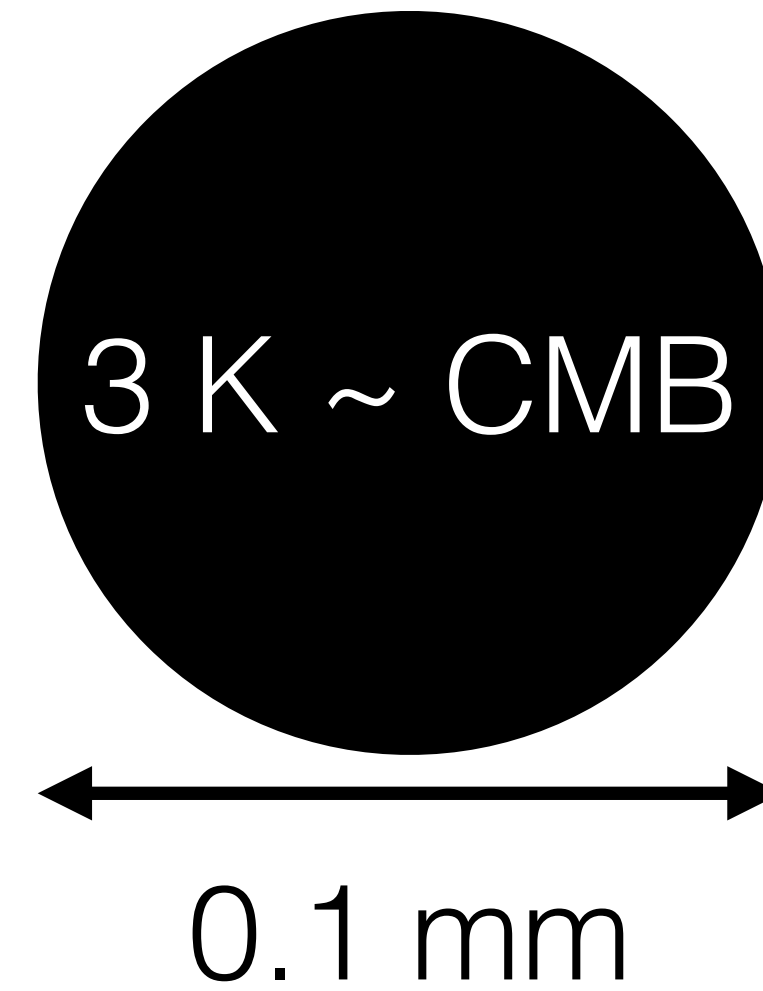
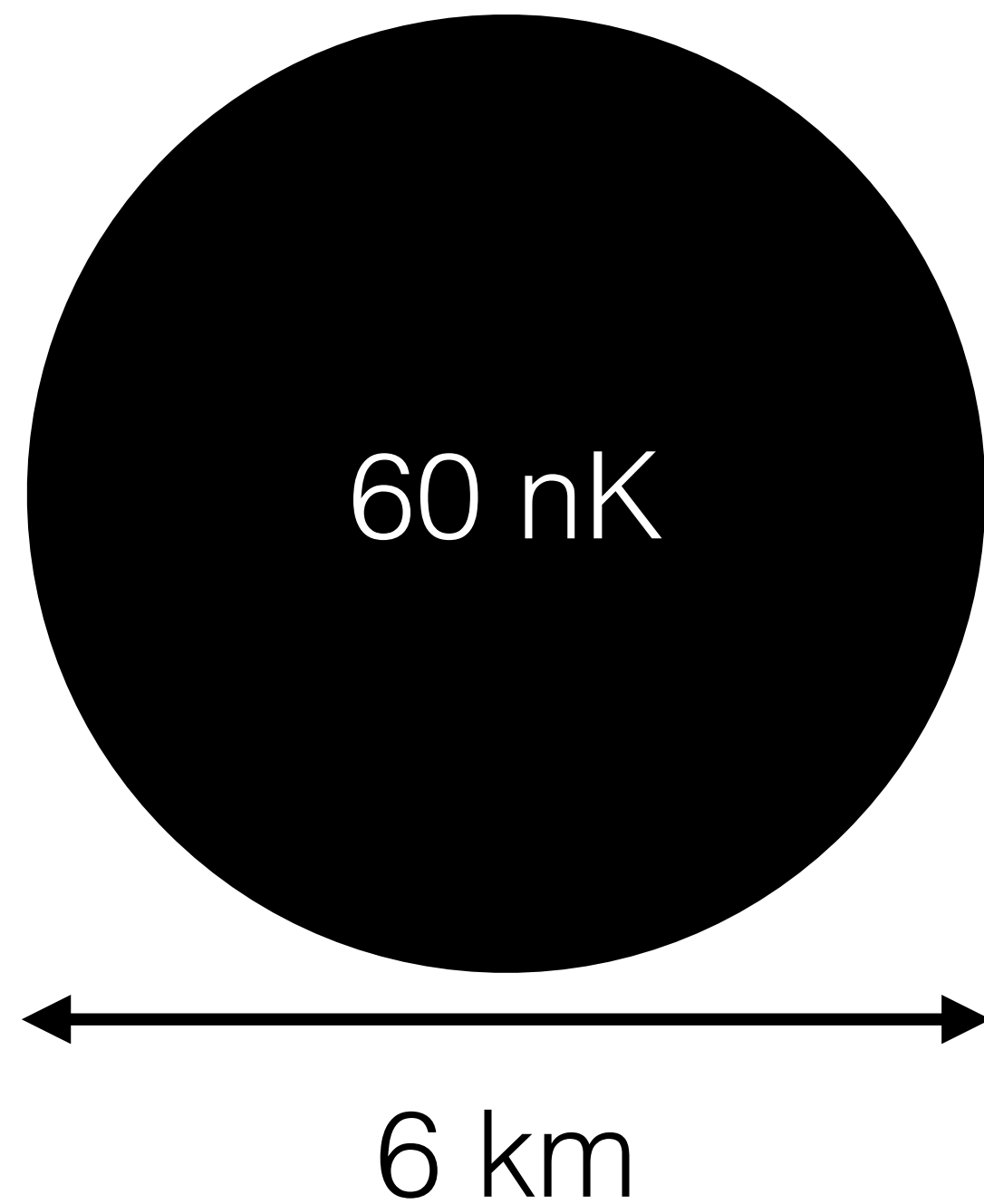
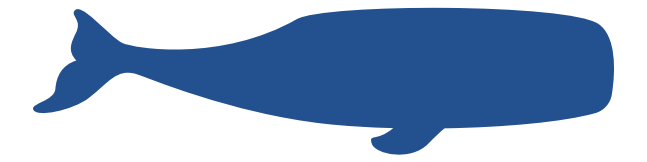
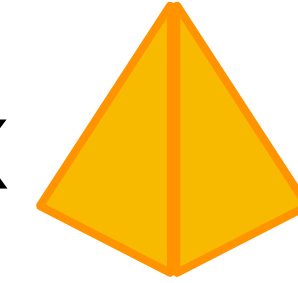
Hawking Radiation

Mass:

Sun

Moon

10 x



Lifetime:

10^{67} years

10^{44} years

13.8 Gyr

1 s

Primordial black Holes



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 - Seeds for SMBHs or Large Scale Structure
 - Could possibly constitute dark matter



Mapping the Northern Sky in High-Energy Gamma Rays

HAWC Observatory

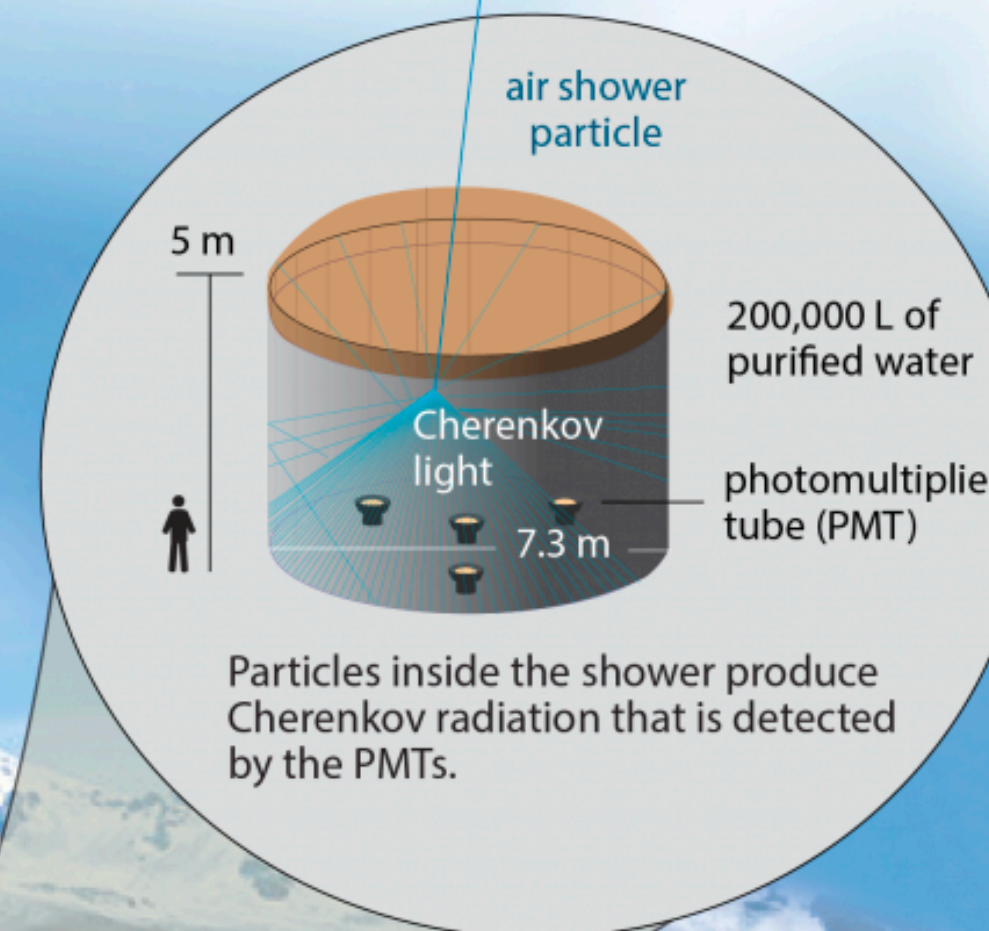
HAWC operates day and night, providing a large field of view for the observation of the highest energy gamma rays.



Pico de Orizaba
(5,626 m)

Water Cherenkov tank

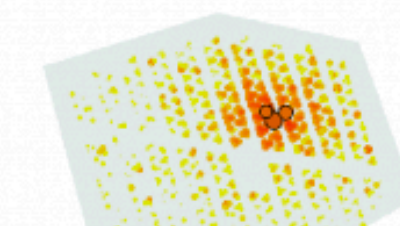
HAWC comprises an array of 300 tanks that record the particles created in gamma-ray and cosmic-ray showers.



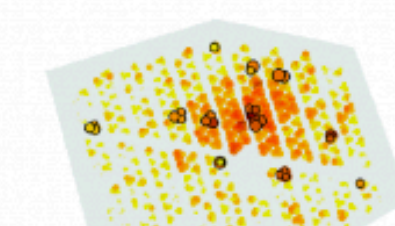
Gamma rays vs cosmic rays

HAWC selects gamma rays from among a much more abundant background of cosmic rays.

gamma-ray shower



cosmic-ray shower



HAWC is located at 4,100 m above sea level, covering an area of 20,000 m².

150 m

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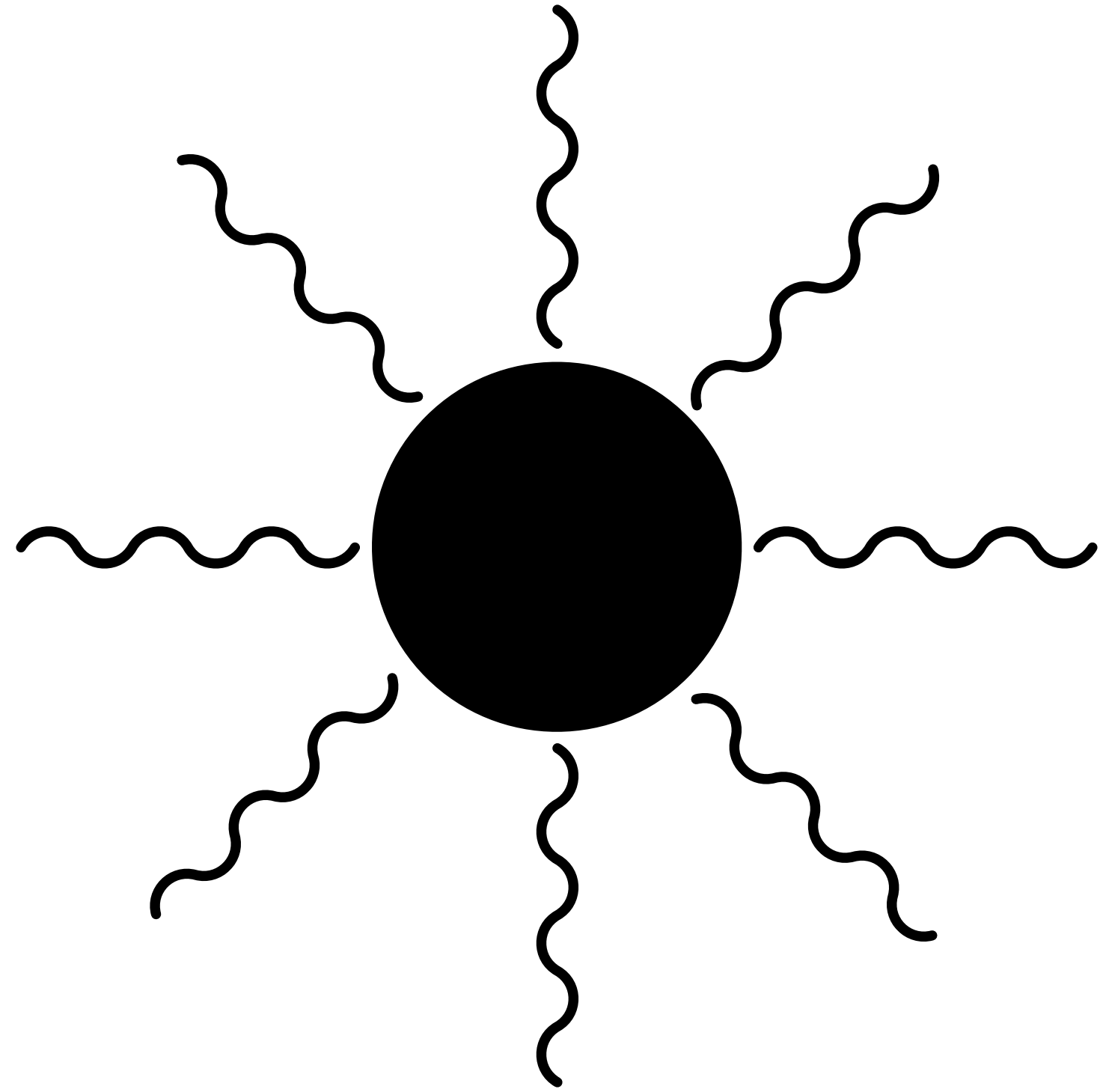
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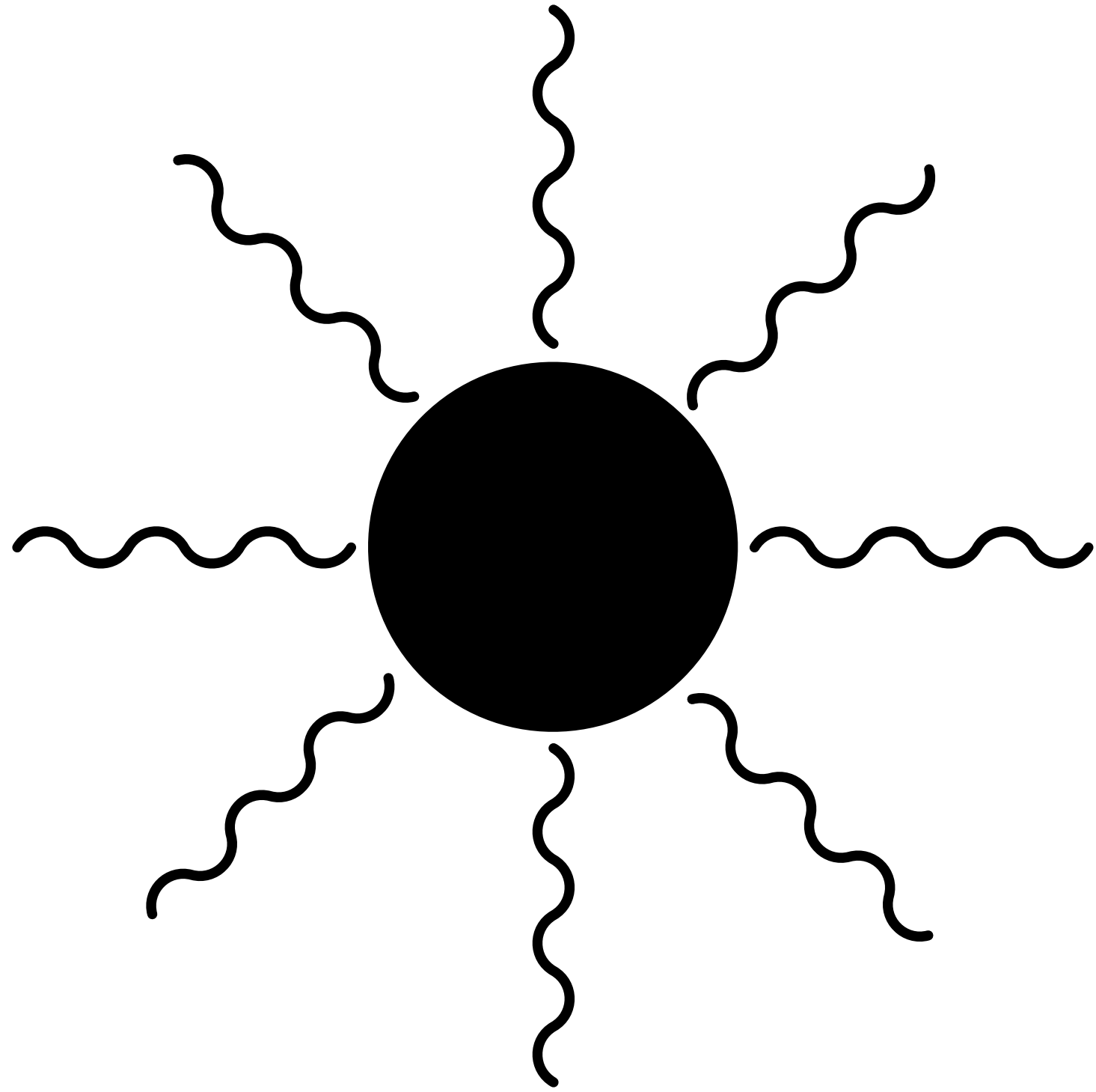
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- We ask the question:
 - What could we learn about BSM physics (e.g., dark sectors) if an evaporating black hole were observed today?

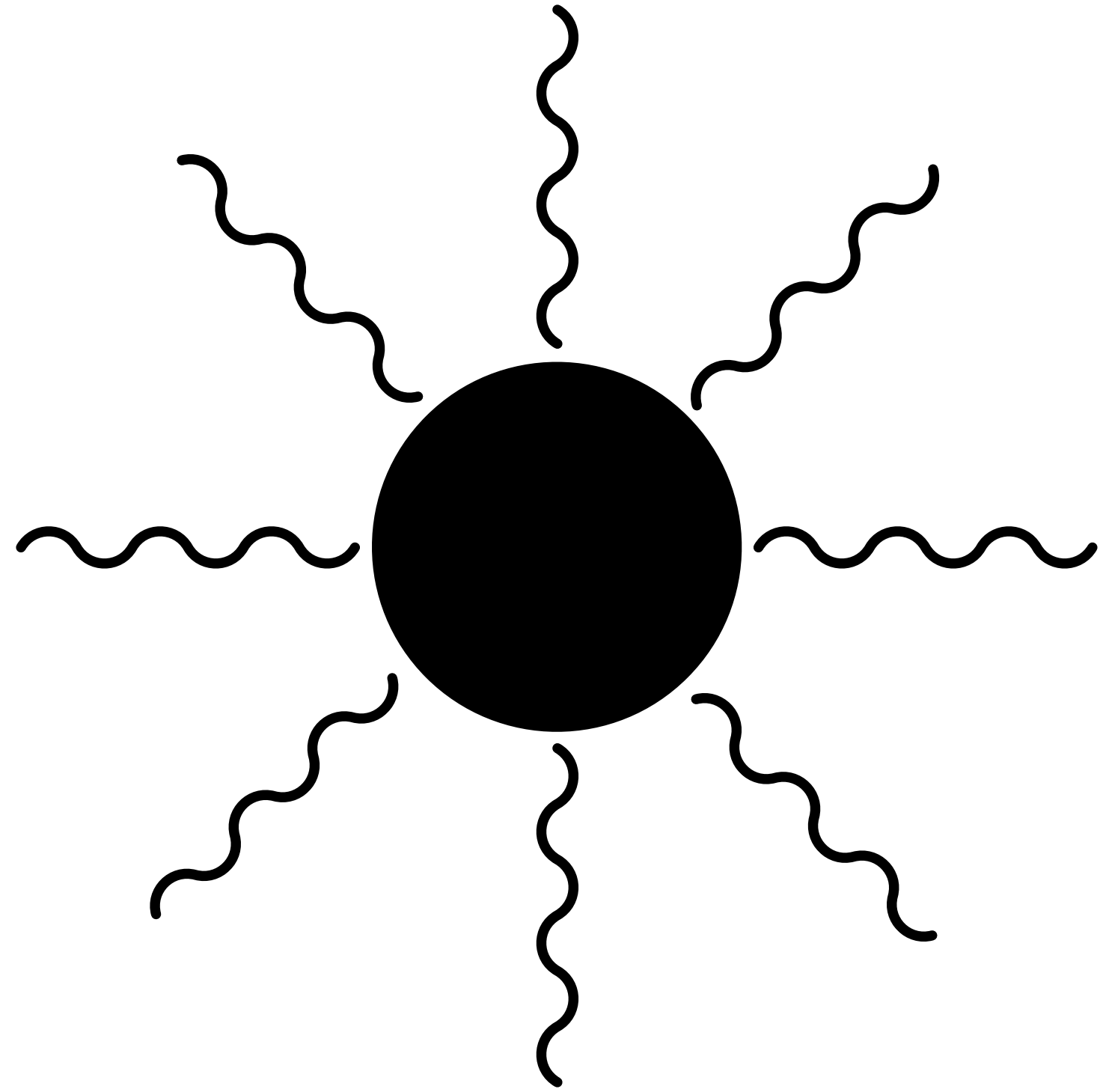
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$$\frac{d^2 N_p^i}{dt dE} = \frac{n_{\text{dof}}^i \Gamma^i(M, E)}{2\pi (e^{E/T} \pm 1)}$$

$$T = \frac{1}{8\pi GM}$$



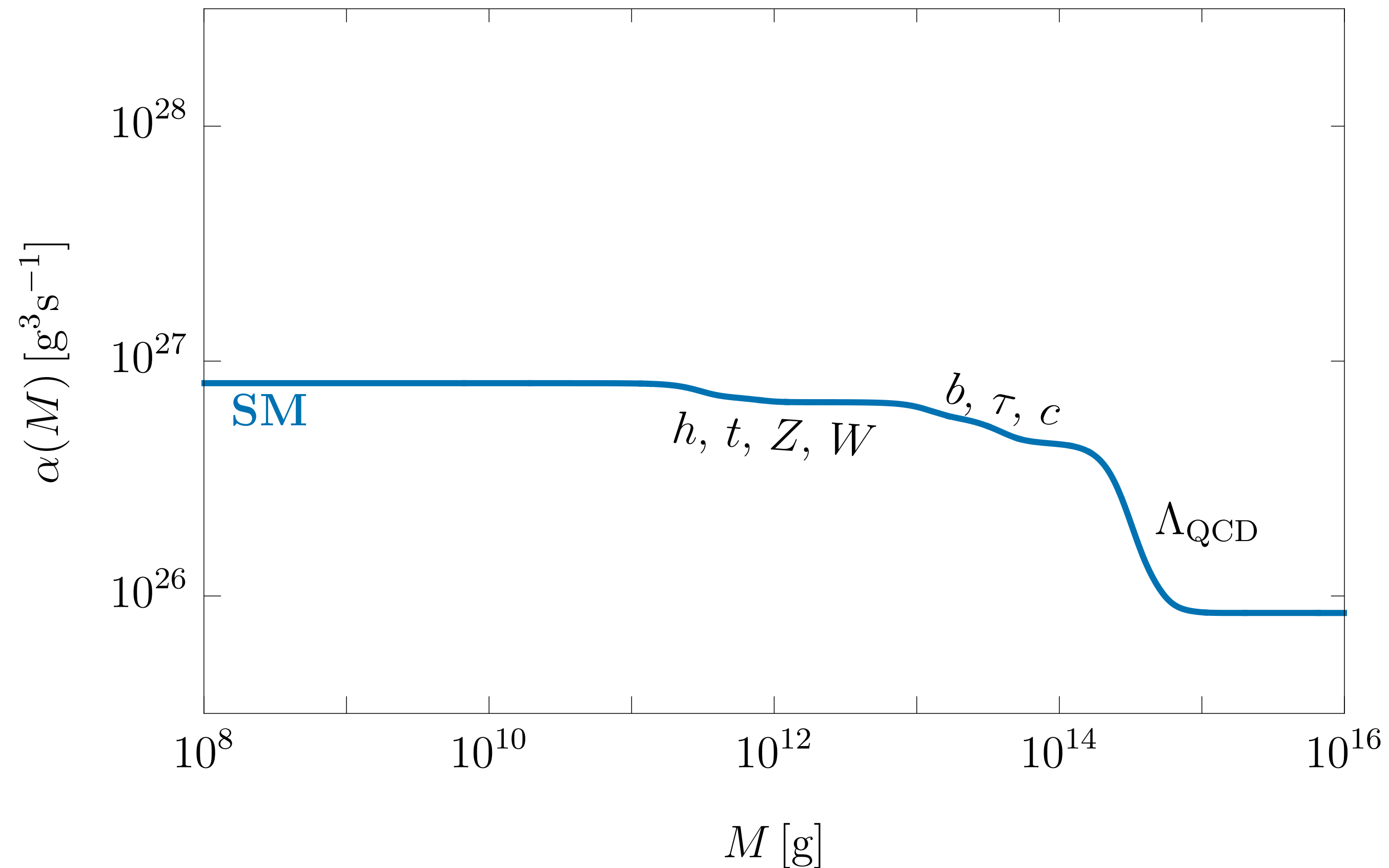
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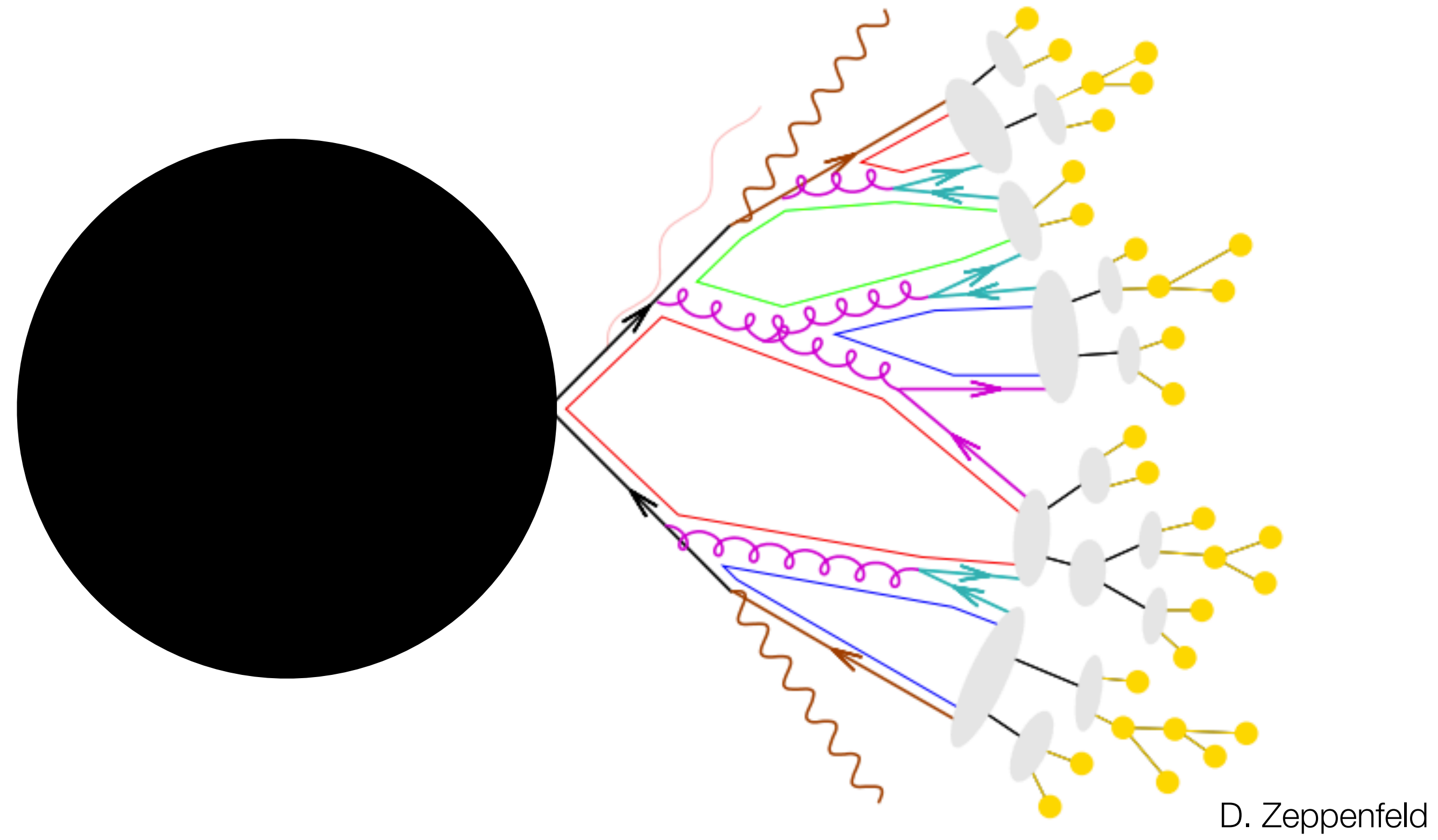
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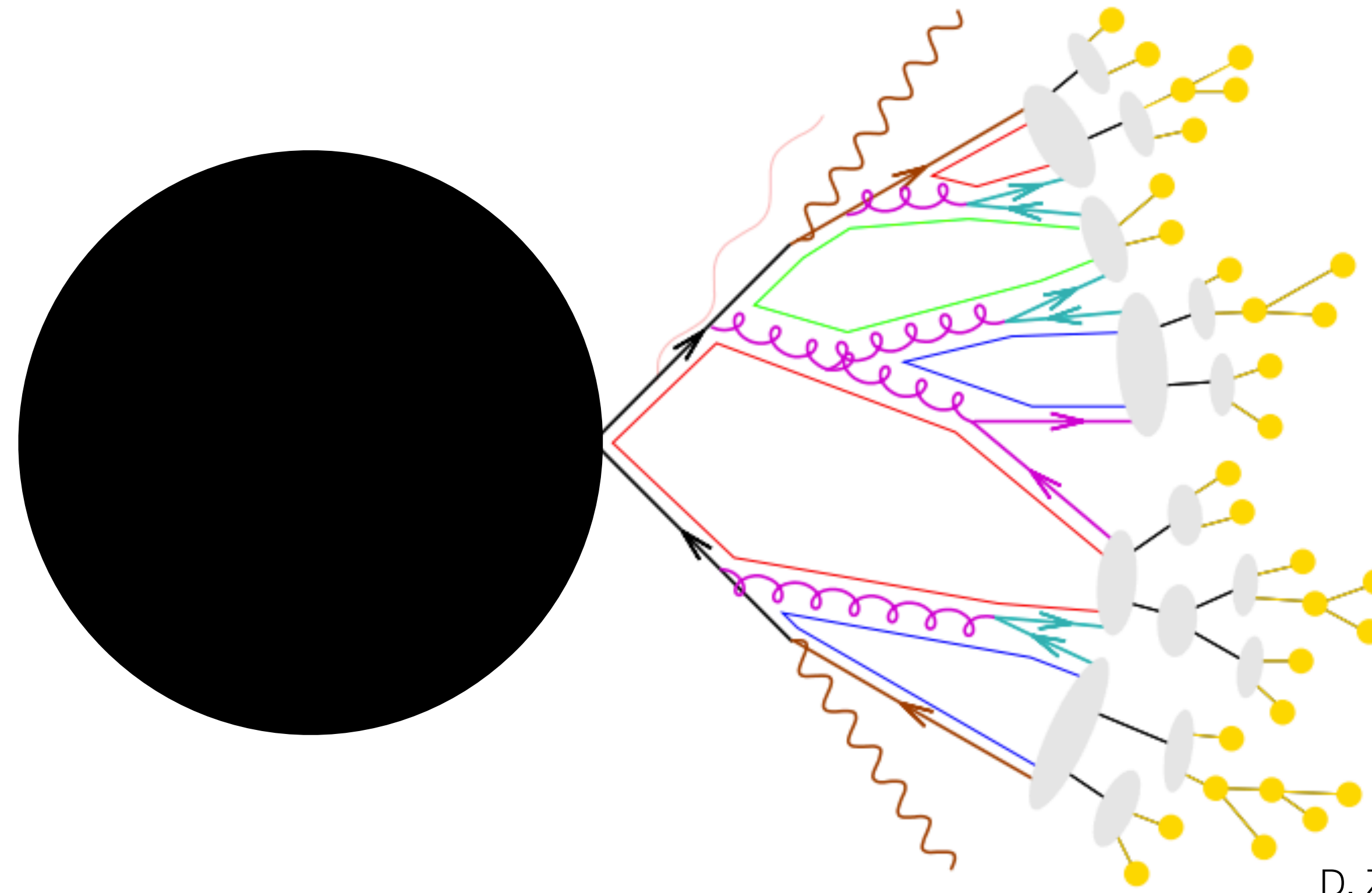
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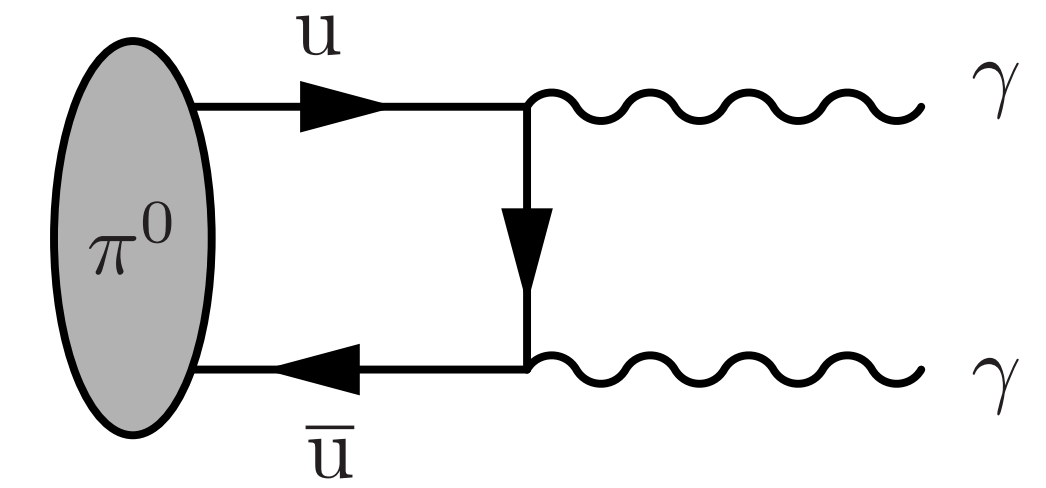
Secondary Photons

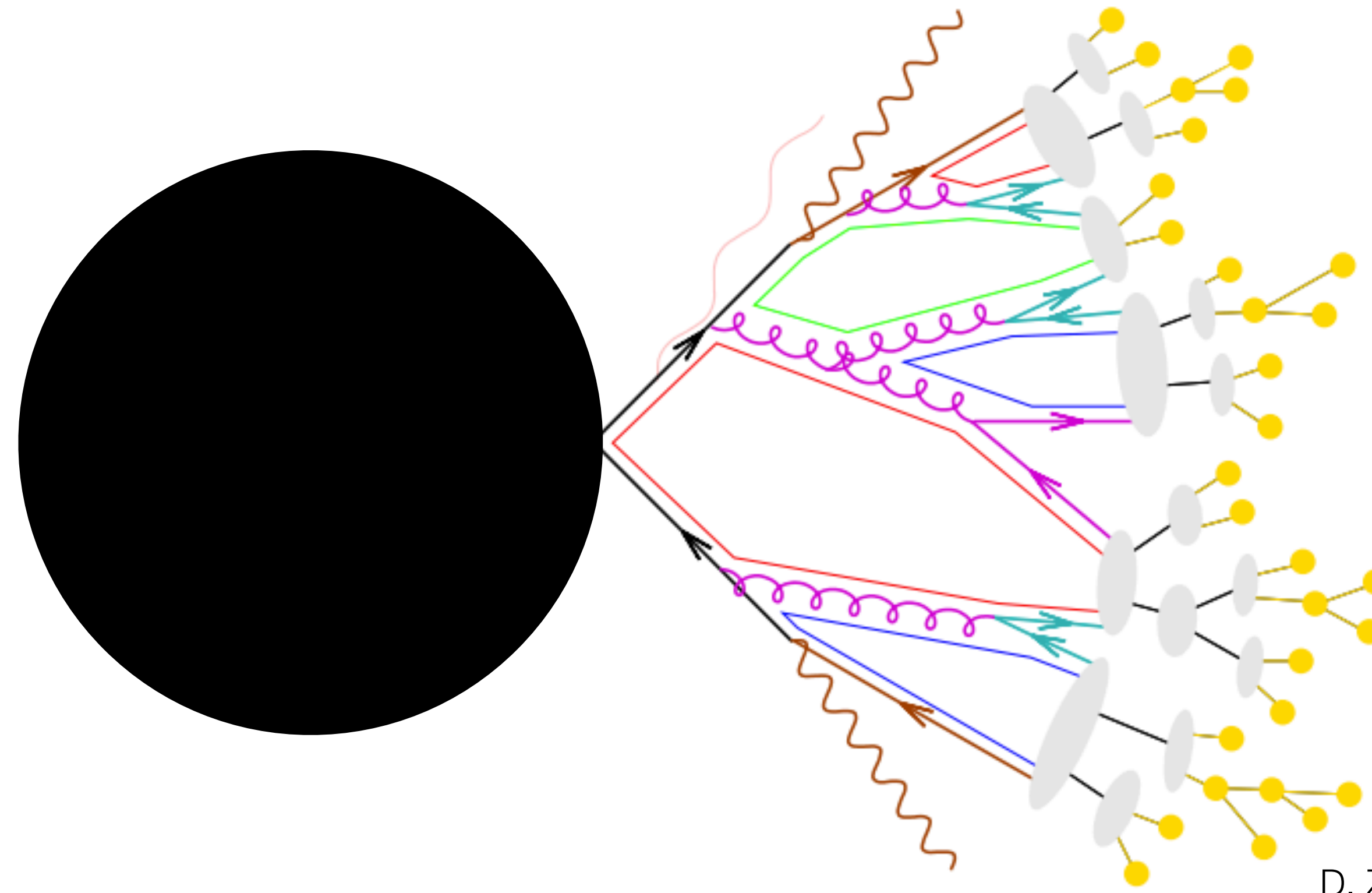


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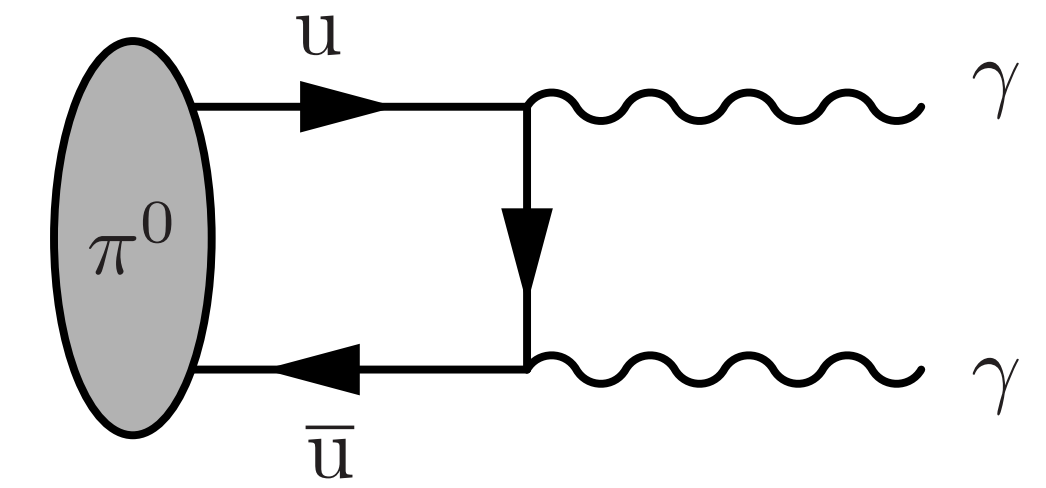


D. Zeppenfeld

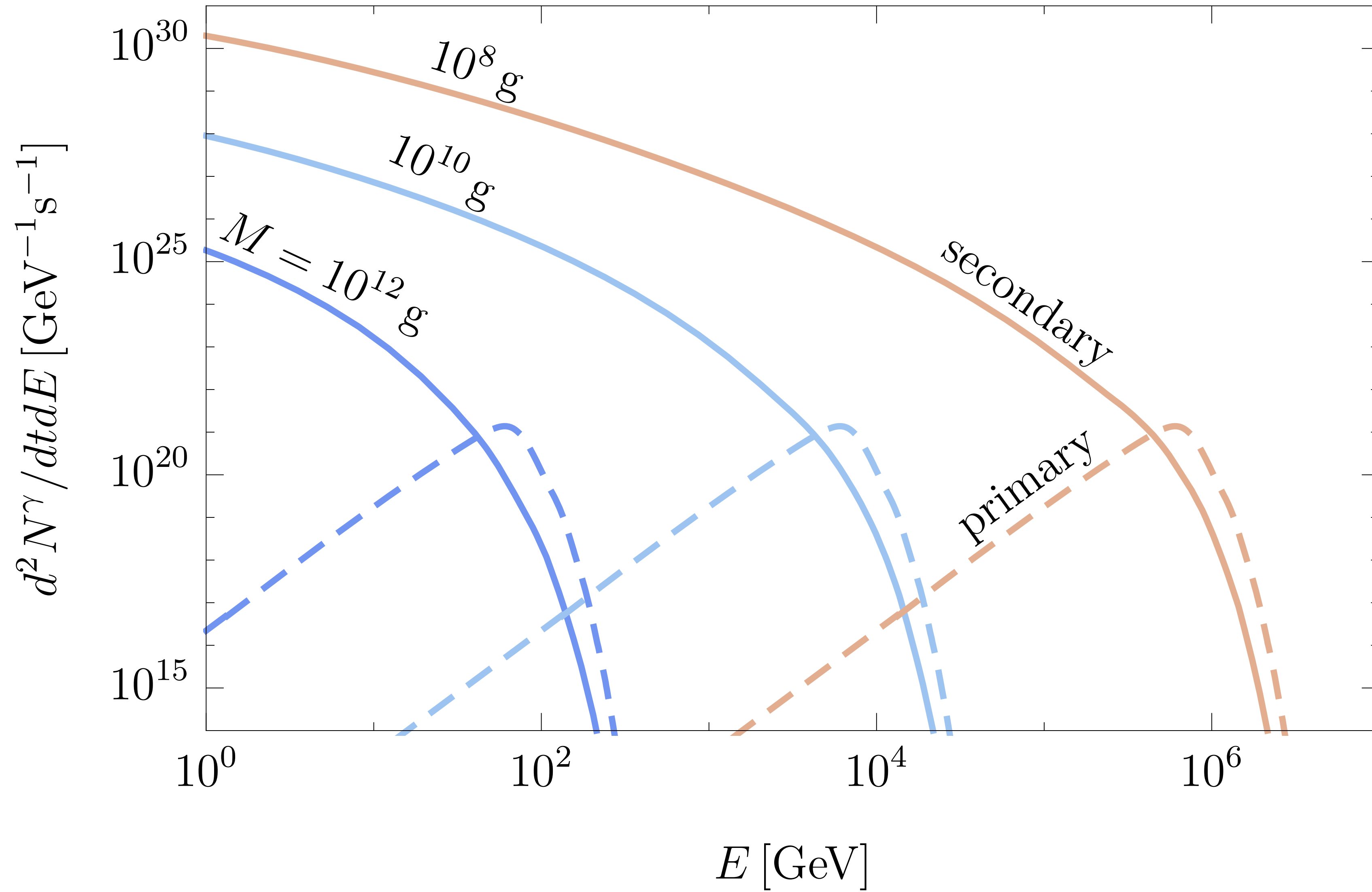




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$$\frac{d^2 N_s^\gamma}{dt dE} = \sum_i \int_0^\infty \frac{d^2 N_p^i}{dt dE_p} (M, E_p) \frac{dN^{i \rightarrow \gamma}}{dE} (E_p, E) dE_p$$



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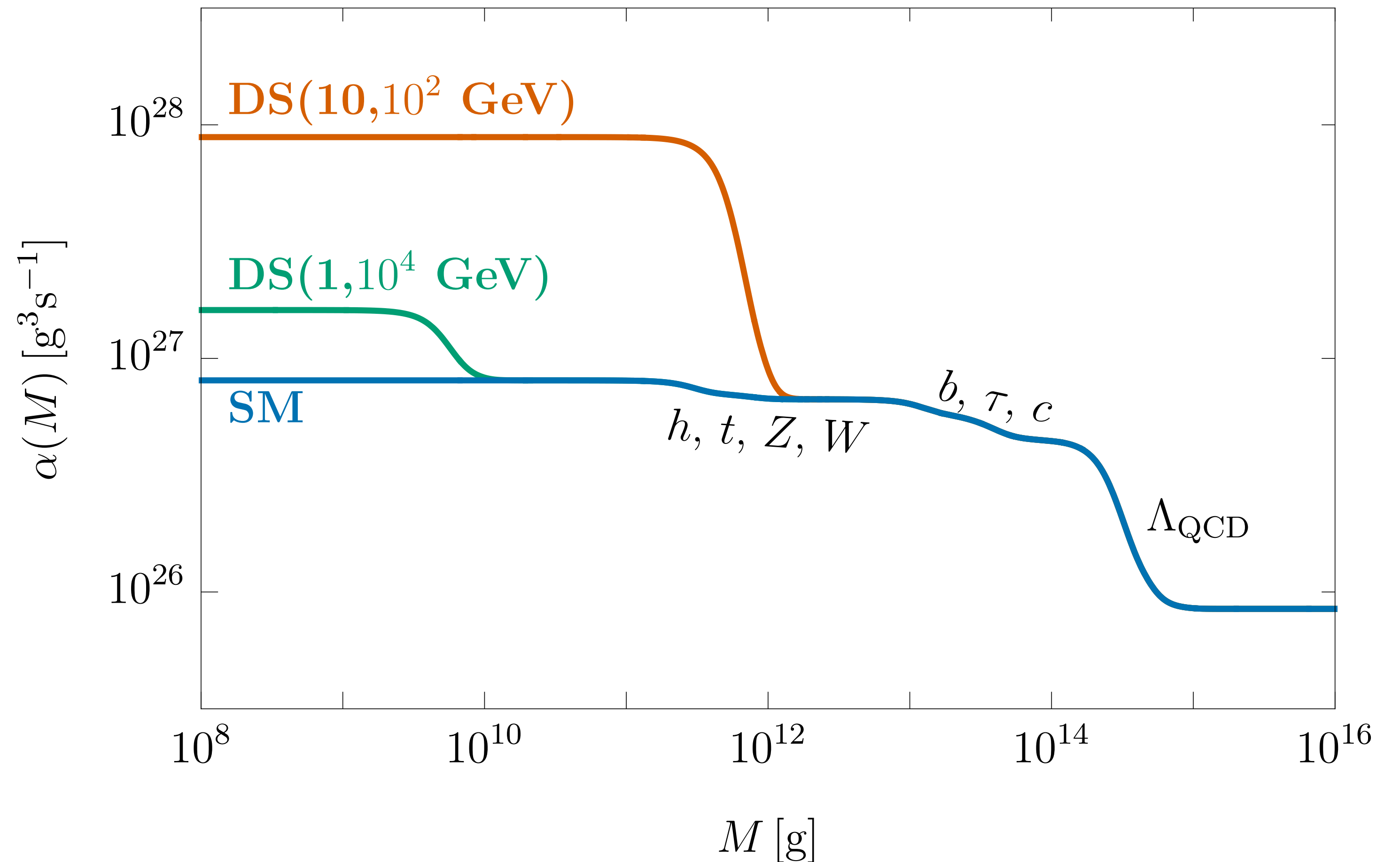
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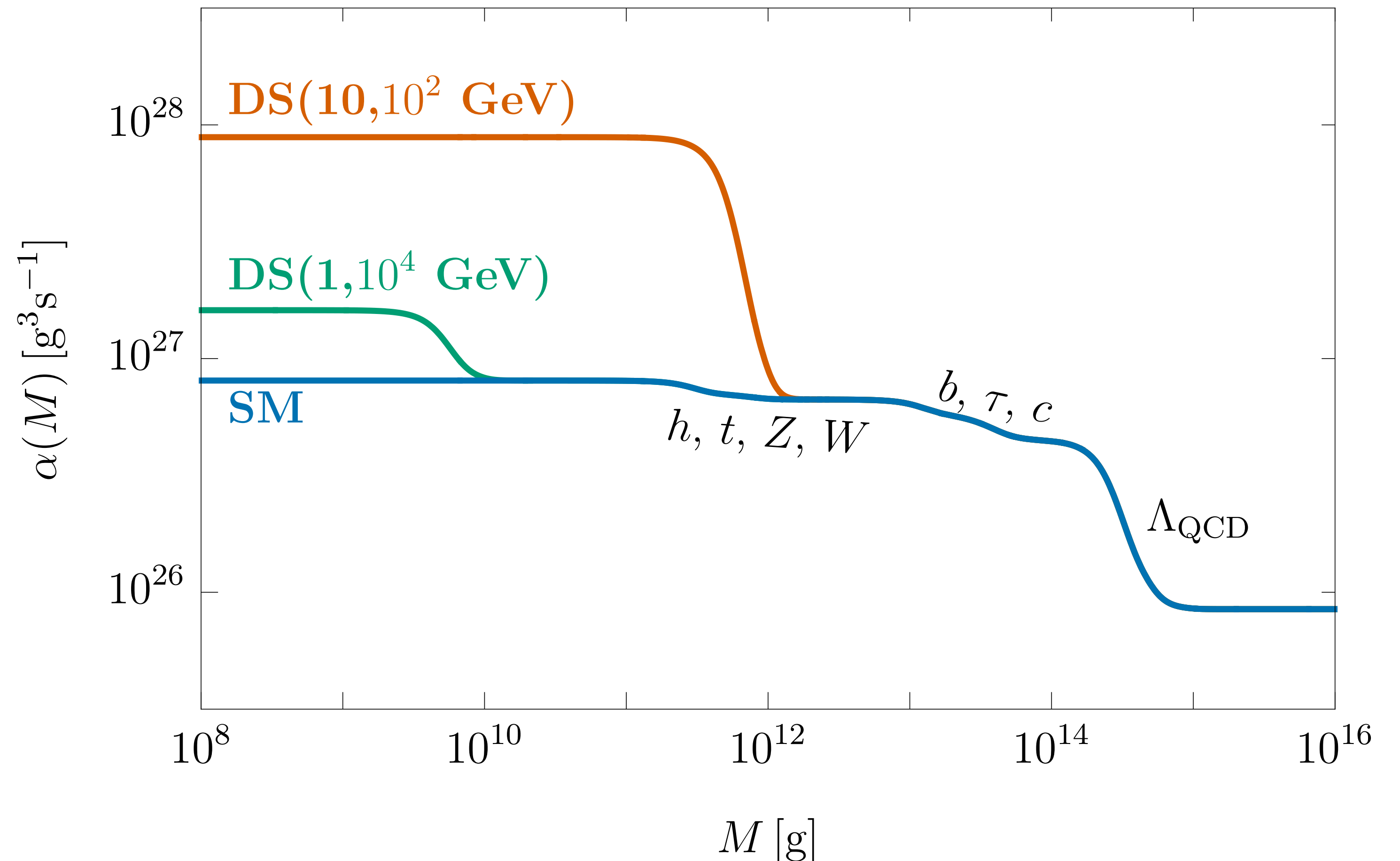


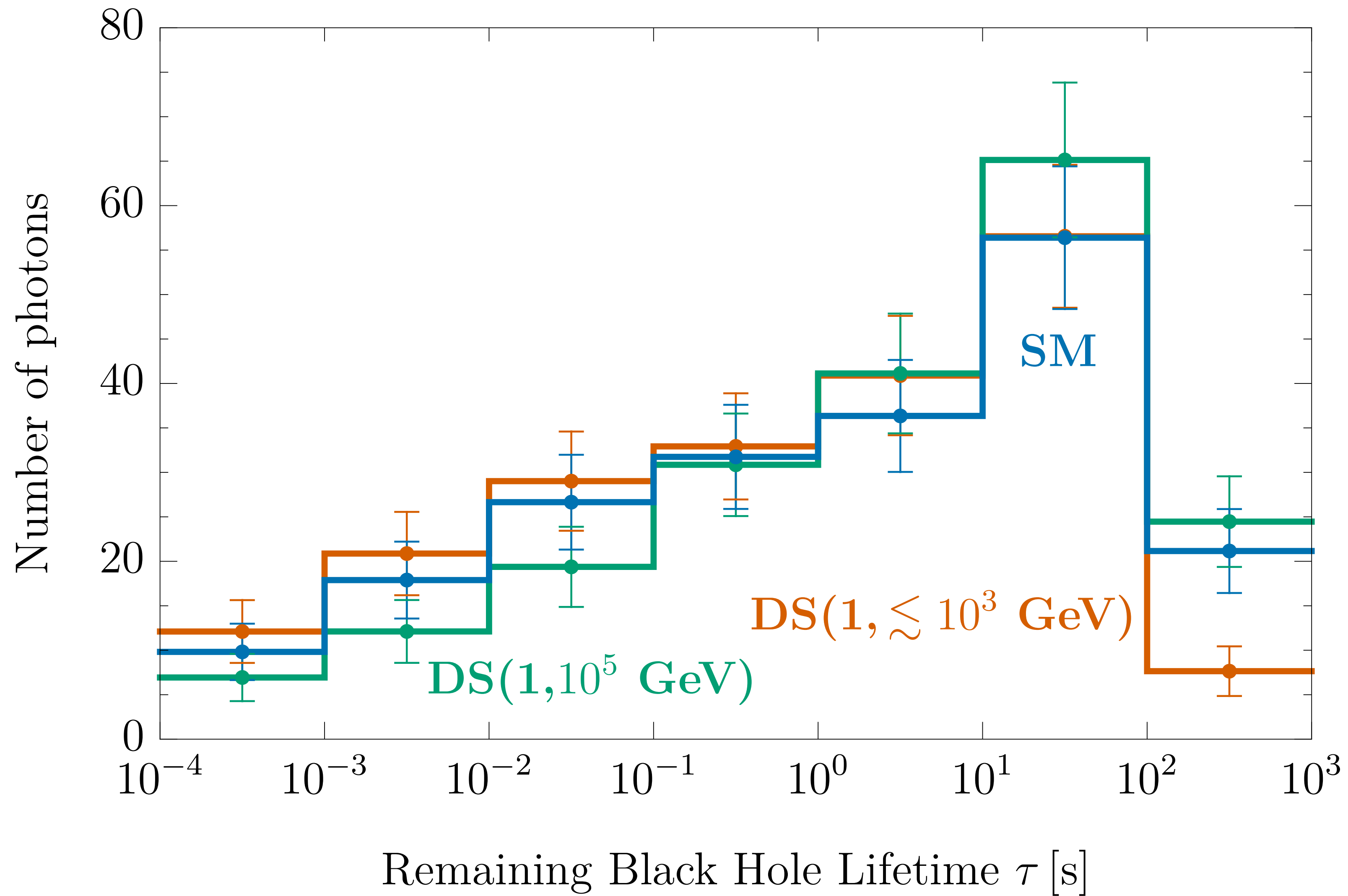
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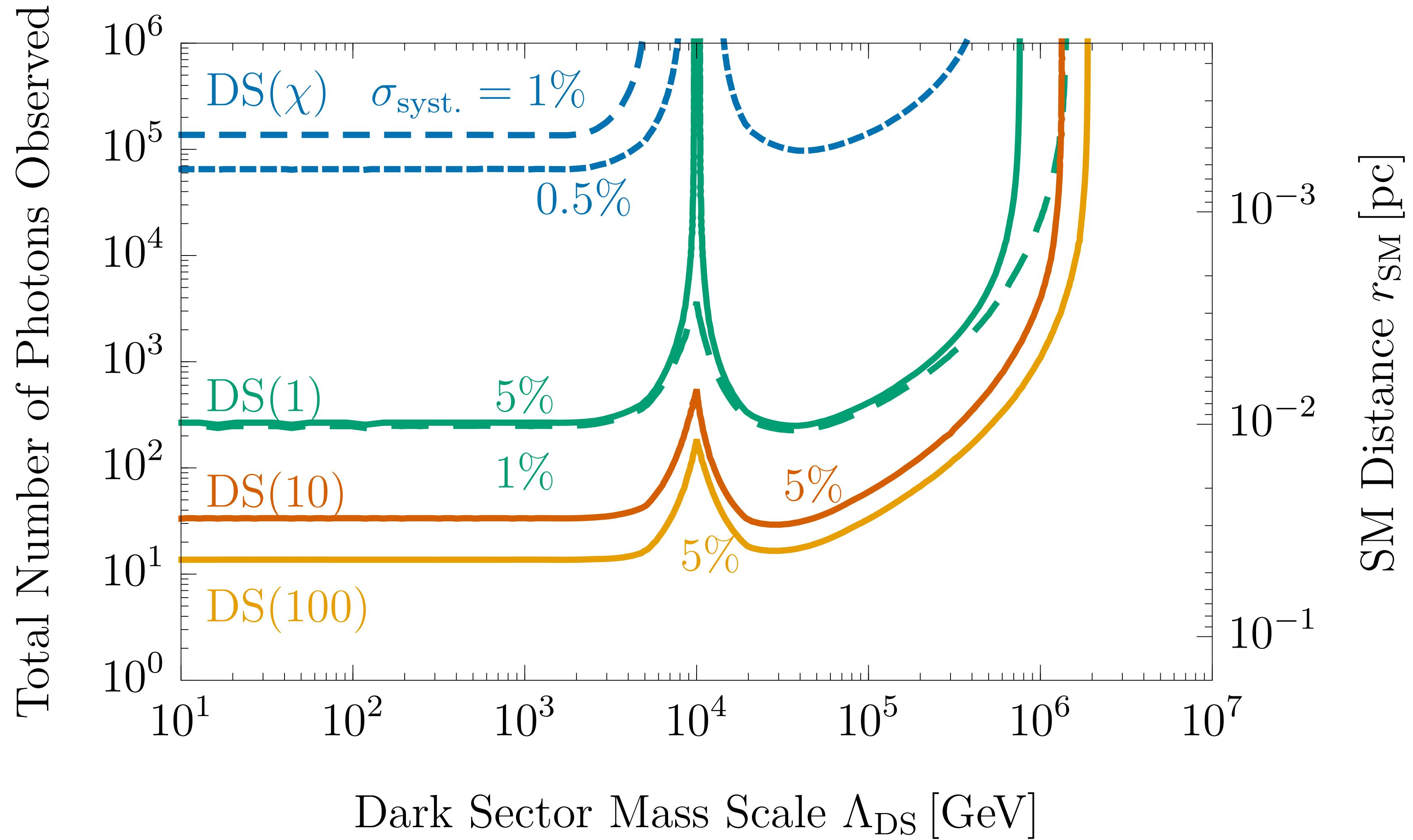
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DS consists of N copies of the SM particles at common mass scale Λ_{DS}







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Table 1: PBH Limits vary with distance scales: (1) from 100 MeV extragalactic γ -ray background assuming no clustering [1,6], (2) from 100 MeV anisotropy measurement [7], (3) from antiproton flux [8] and (4) from Very High Energy (VHE) searches [9].

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Production & propagation of antiprotons, QCD, BSM

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