Dark Matter and Dark Sector 2: Detection

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Recap: lecture 1

- We have much astrophysical evidence that dark matter exists

- We do not know what it is, but we know what properties it must have and non of the SM particles fit the profile

- We know that: 1. It must be stable
  2. Non relativistic
  3. Must not interact via SM charges

- Many possibilities of what DM could be

- Very well motivated possibility called WIMP

This lecture we will talk about how to detect dark matter

Late 1970s - Vera Rubin

Image: Carnegie Institute for Science

First scientist to measure star speeds with very high accuracy

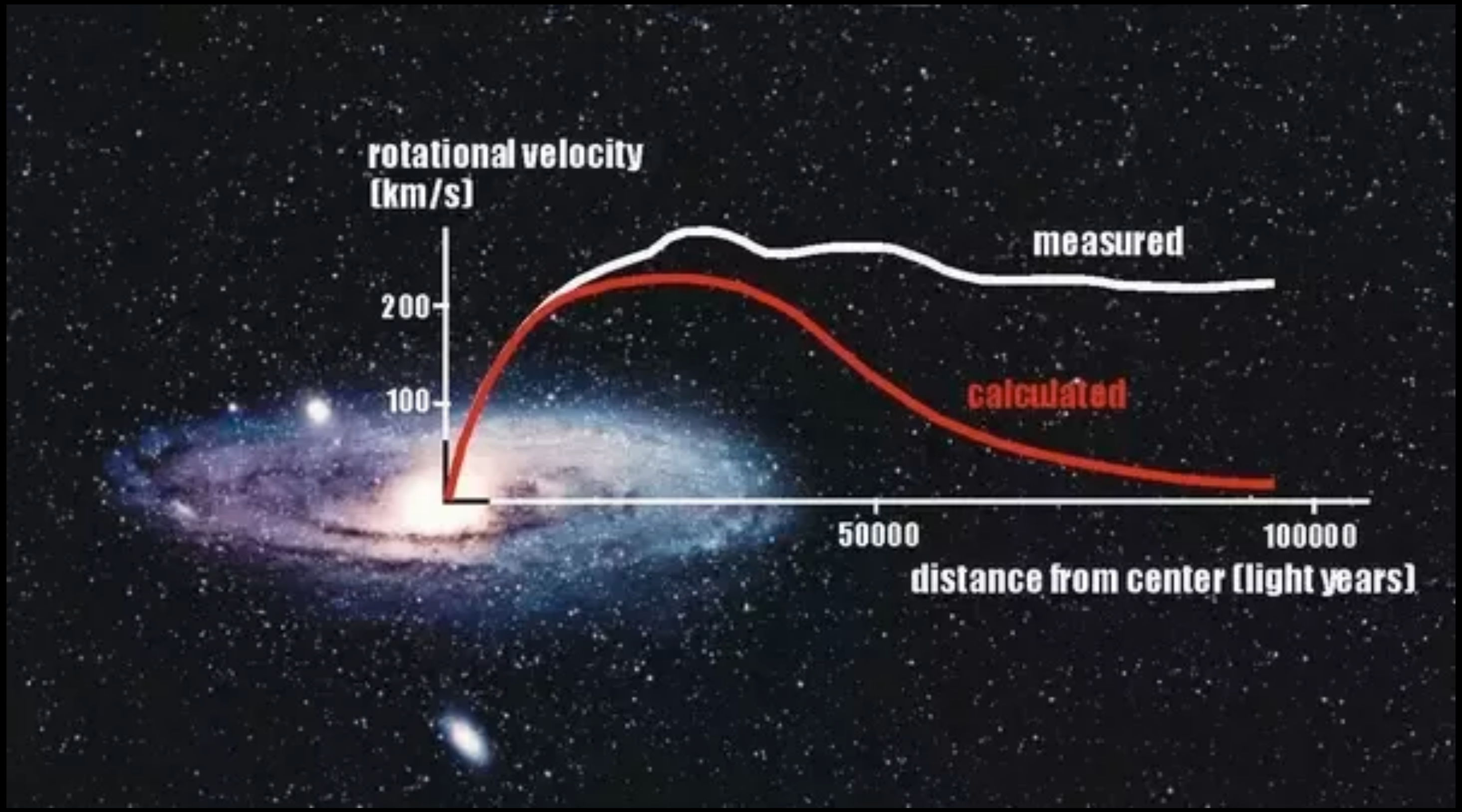


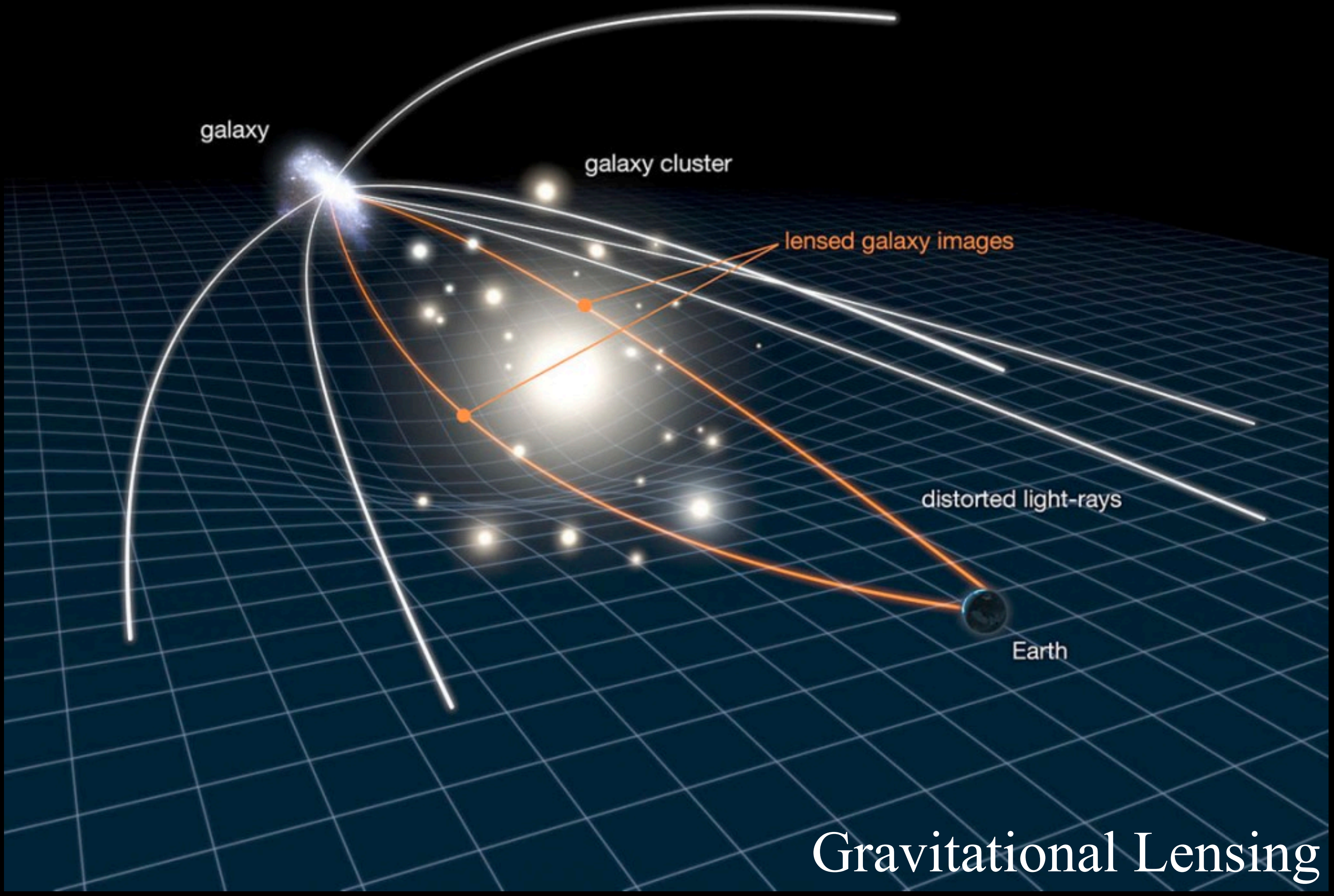
Image: quora.com

$$\frac{mv^2}{r} = \frac{GMm}{r^2}$$

$$v = \sqrt{\frac{GM}{r}}$$

Lets look at scales larger than galaxies

Again, lets imagine the galaxy cluster is invisible



We still see this bending of light

What do galaxy clusters tell us about DM?

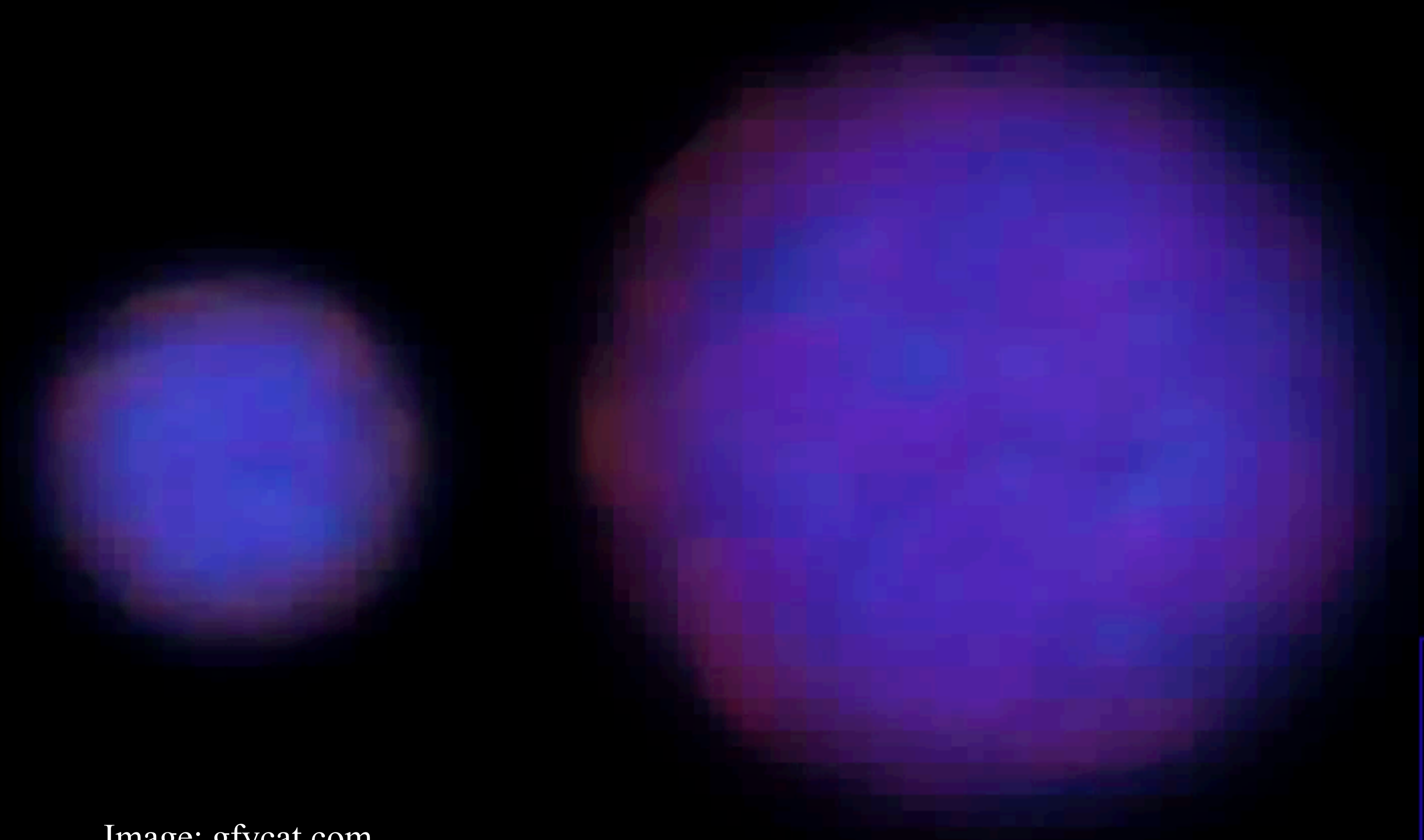


Image: gfycat.com

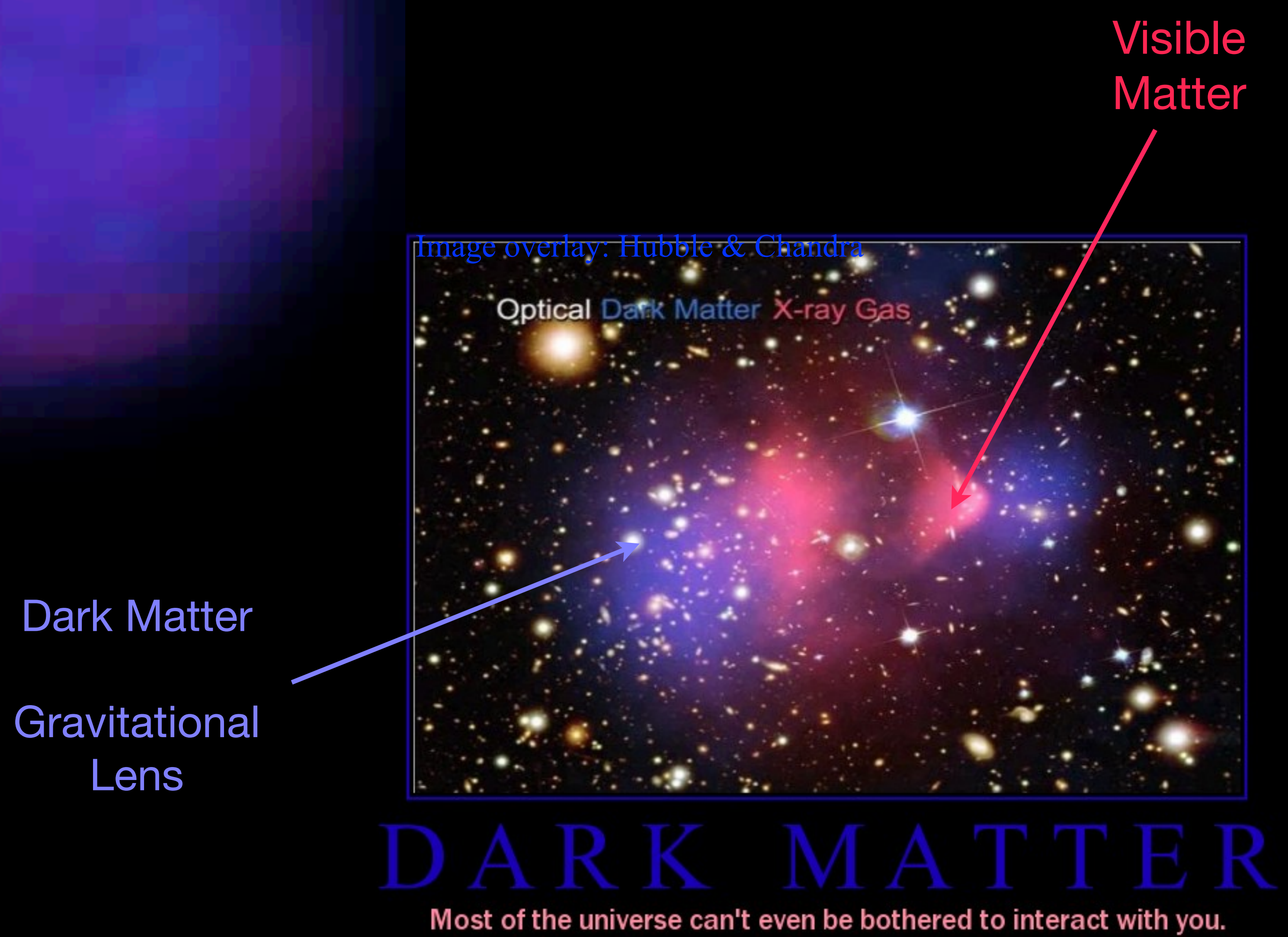


Image: yumpu.com

What do cosmological observations tell us?

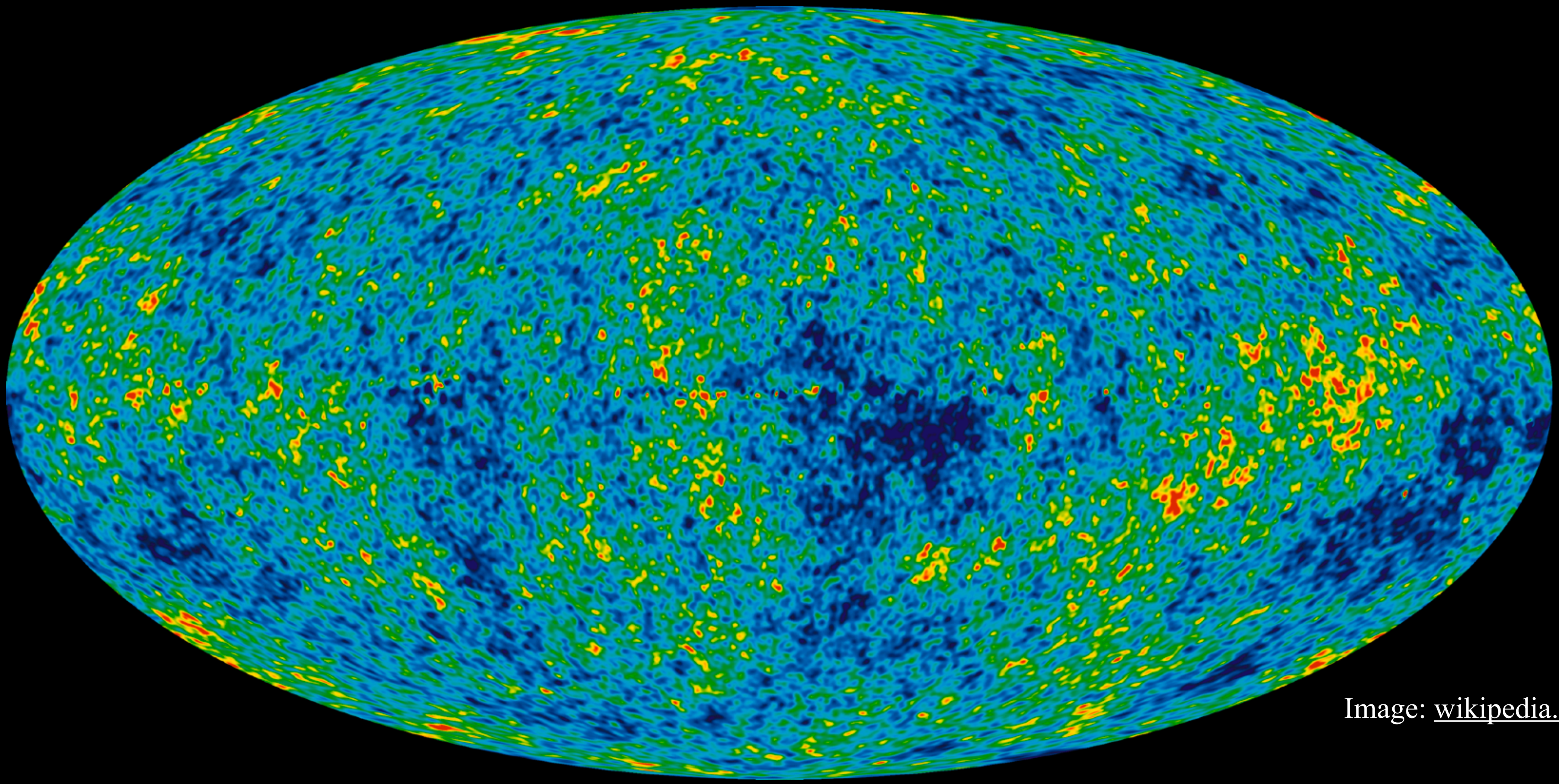


Image: wikipedia.org

Epoch of recombination

- Neutral atoms were formed, photons could move freely since they were no longer locked to charged particles

- These free moving photons reach us today

What do cosmological observations tell us?

CMB Power Spectrum -
gives cosmologists a way to mathematically understand fluctuations

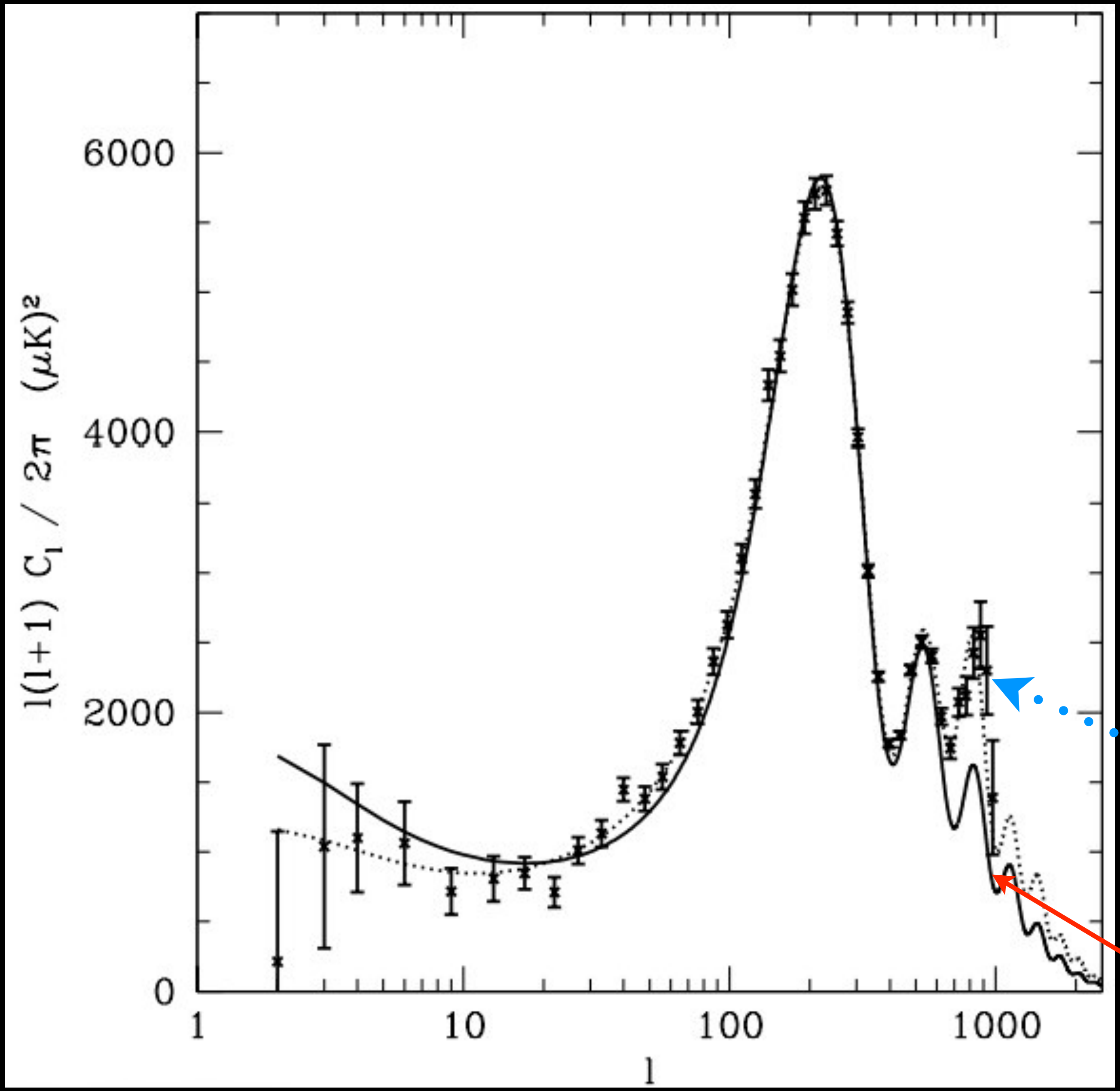


Image: Sean Carrol

Dark Matter

MOND -
Modification of Gravity

Evidence shows dark matter exists at largest scales

What is Dark Matter made of?

We simply have no idea.

We DO know:

- It must be cold (non-relativistic) at the time of structure formation
- It must be super long-lived or completely stable
- It must be some new state lying beyond the SM

Non-EM interacting

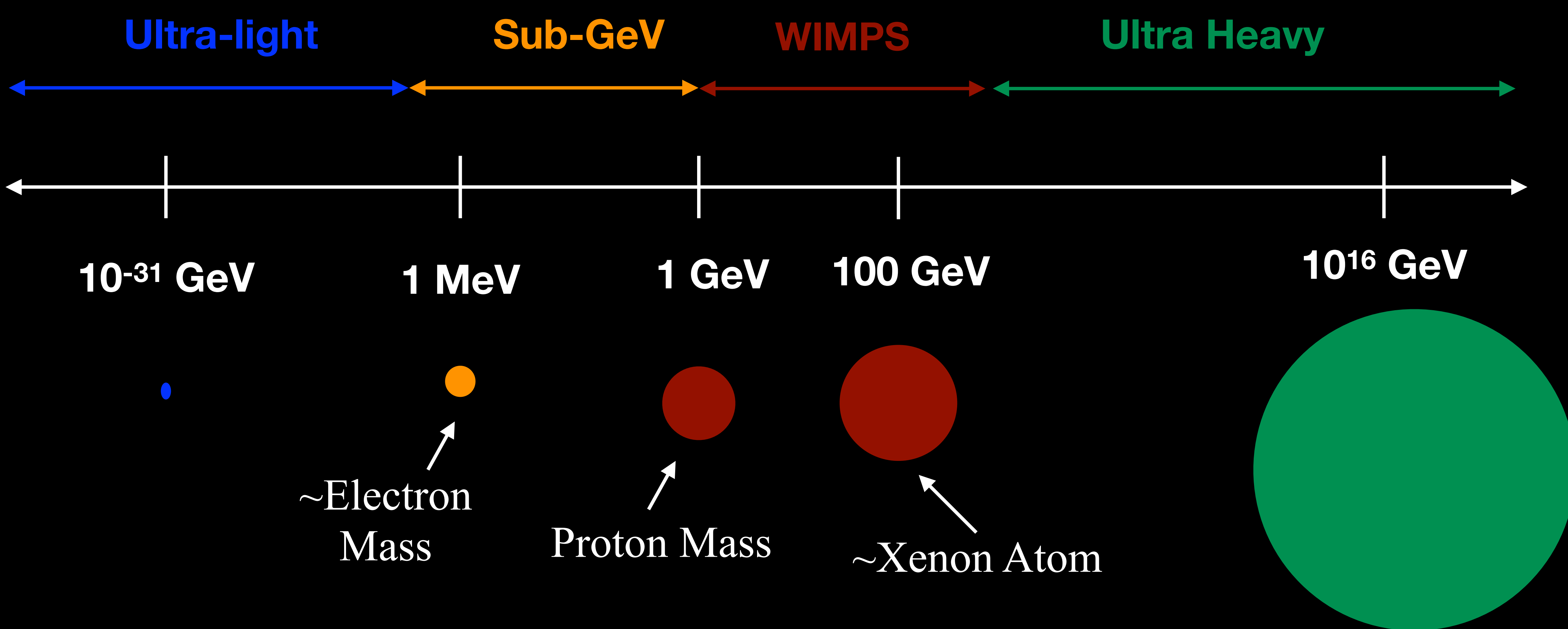
Non-QCD interacting

Dark Matter should be described by a quantum field corresponding to a definite spin, uncharged under $U(1)_{EM}$ or $SU(3)_{C}$.

(So: no tree-level interactions with gluons or photons).

- It **may** interact with the SM through some new force

Range of possibilities is VAST



Primordial Blackholes - much heavier than Ultra heavy

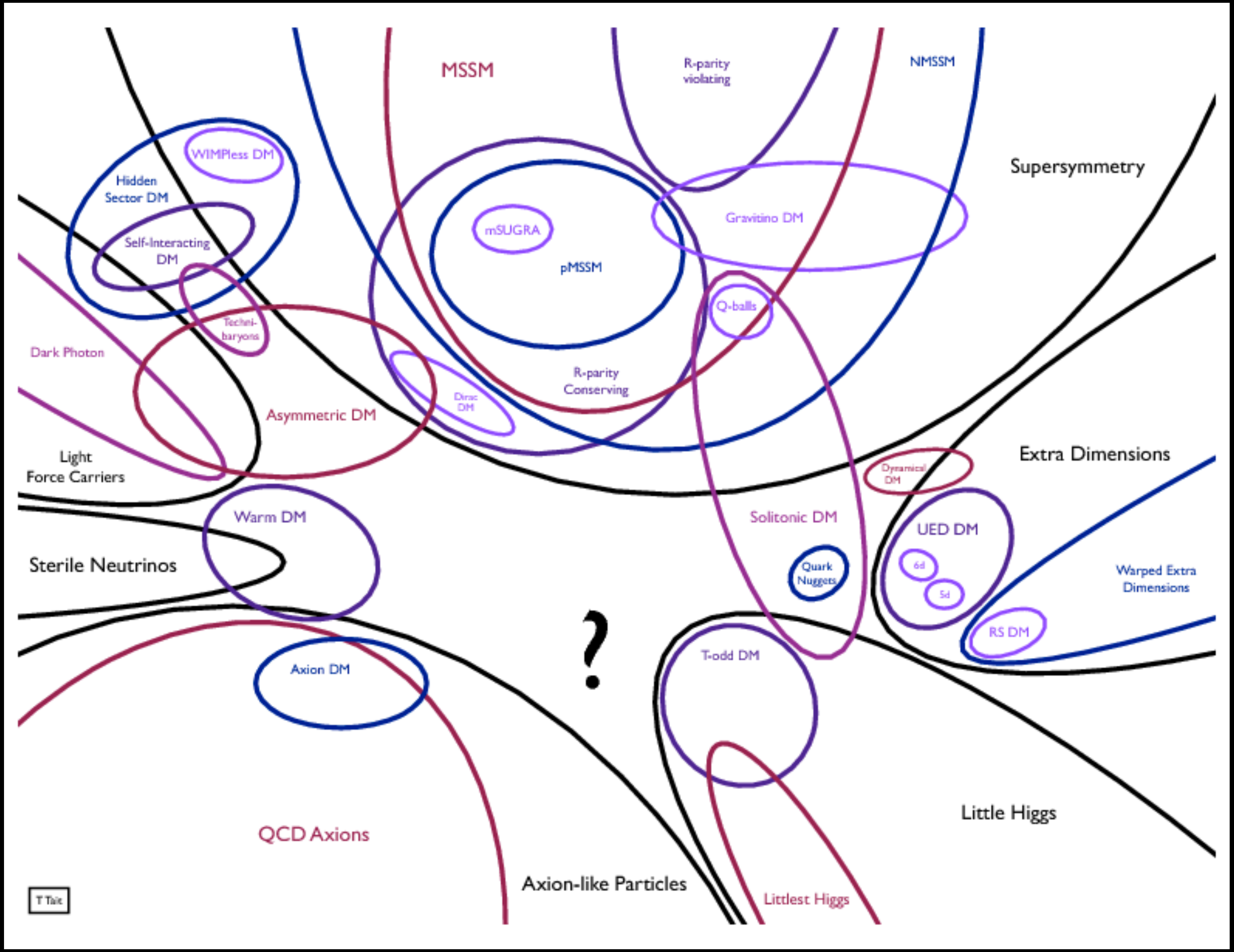
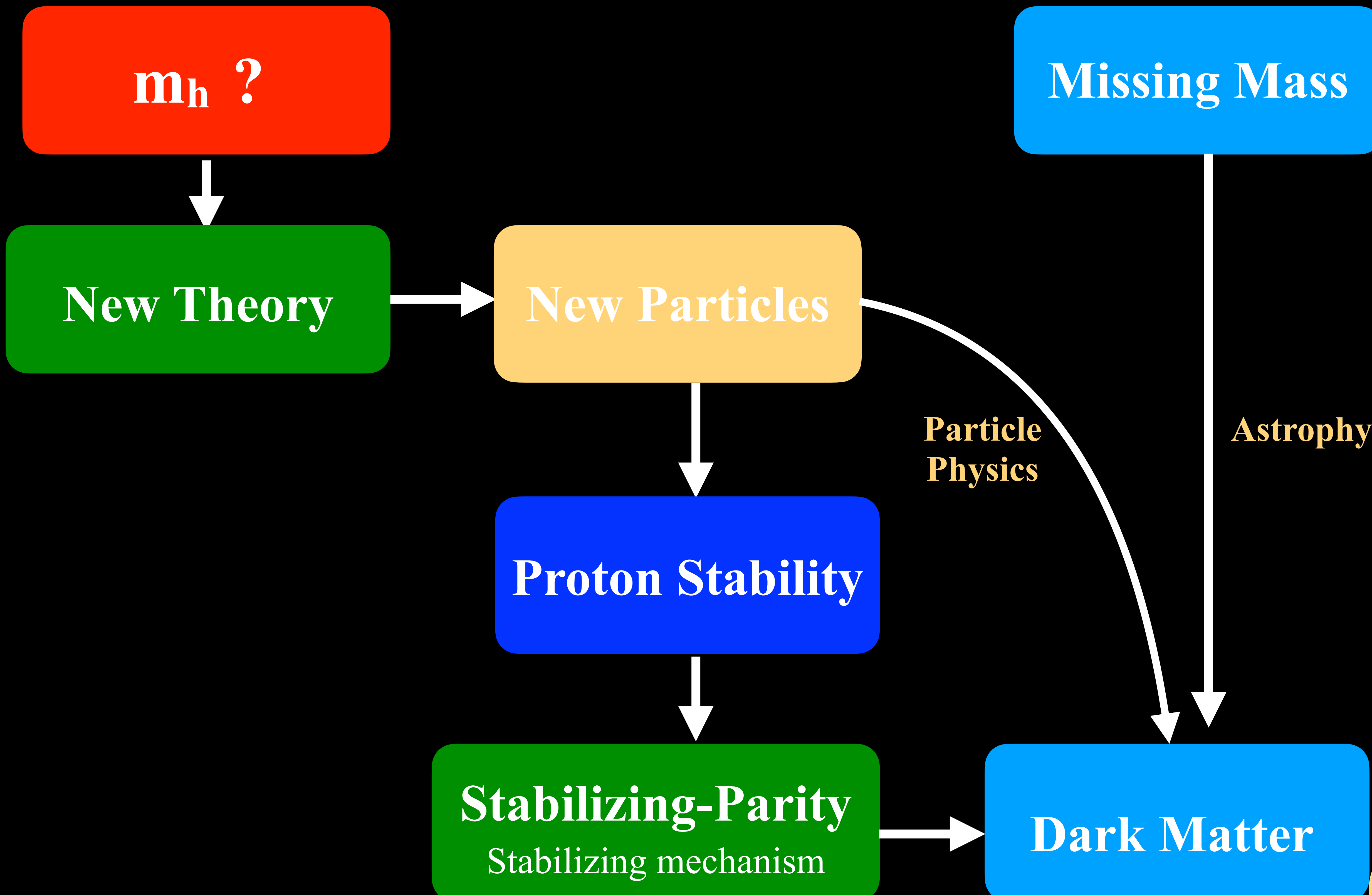


Image: ps.uci.edu

Range of possibilities is VAST

Solutions of hierarchy problem



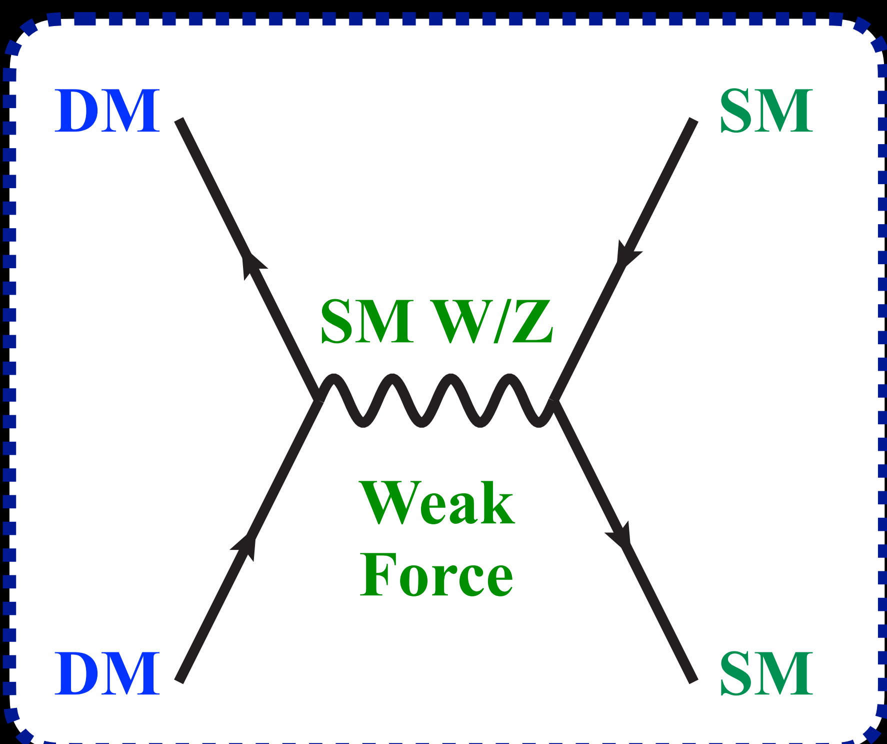
Solving the Hierarchy problem resulted in a perfect class of dark matter candidates called

Weakly-Interacting Massive Particles

"Electroweak" interactions ($W^{\pm}$, Z, h)

"weak-scale" mass
1 - 10 000 GeV

Explains: **Why so much dark matter around**



Dark matter annihilation

Observed **amount of dark matter** today → 'Relic Density'

$$\Omega h^2 \sim \frac{0.1\ \mathrm{pb}}{\langle \sigma v \rangle}$$

Weak scale
annihilation rate

In this lecture we will cover

1. How is dark matter produced in the early universe?

   How do we get its relic abundance?

2. How do we search for it today?

   In the sky (Indirect detection)

   In accelerator experiments

   Underground (Direct detection)

Relic density of dark matter

Expansion rate of the universe described by Hubble rate H(z)

In the standard model of cosmology called Lambda Cold Dark Matter (ΛCDM), Hubble rate is

$$H(z) = H_0\sqrt{\Omega_m(1+z)^3 + \Omega_{rad}(1+z)^4 + \Omega_\Lambda}$$

Redshift- gives us an idea of cosmological time

Relic density of dark matter

Expansion rate of the universe described by Hubble rate H(z)

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$$H(z) = H_0\sqrt{\Omega_m(1+z)^3 + \Omega_{rad}(1+z)^4 + \Omega_\Lambda}$$

Hubble constant - standard constant to quantify universe expansion

Relic density of dark matter

Expansion rate of the universe described by Hubble rate H(z)

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$$H(z) = H_0\sqrt{\Omega_m(1+z)^3 + \Omega_{rad}(1+z)^4 + \Omega_\Lambda}$$

Total matter in the universe -
includes SM and dark matter

$$\Omega_m = \Omega_c + \Omega_b$$

Relic density of dark matter

Expansion rate of the universe described by Hubble rate H(z)

In the standard model of cosmology called Lambda Cold Dark Matter (ΛCDM), Hubble rate is

$$H(z) = H_0\sqrt{\Omega_m(1+z)^3 + \Omega_{rad}(1+z)^4 + \Omega_\Lambda}$$

Total amount of radiation in the universe - relativistic free particles like photons, neutrinos

Relic density of dark matter

Expansion rate of the universe described by Hubble rate H(z)

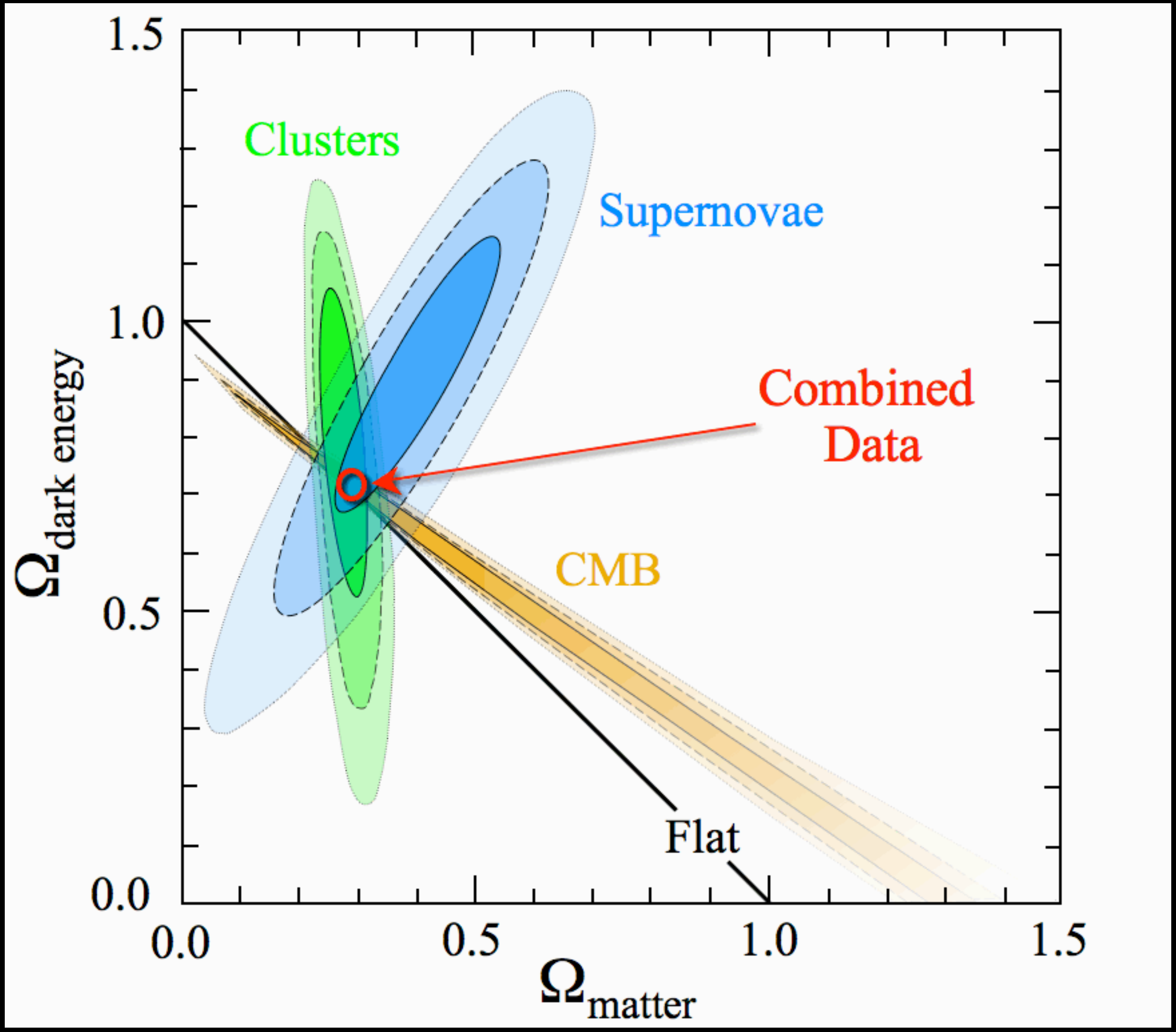
In the standard model of cosmology called Lambda Cold Dark Matter (ΛCDM), Hubble rate is

$$H(z) = H_0\sqrt{\Omega_m(1+z)^3 + \Omega_{rad}(1+z)^4 + \Omega_\Lambda}$$

Dark energy density - phenomenon making universe expand

Relic density of dark matter

Using combined data from a variety of telescopes measuring billions of galaxies, CMB and supernovae, we can fit those parameters.



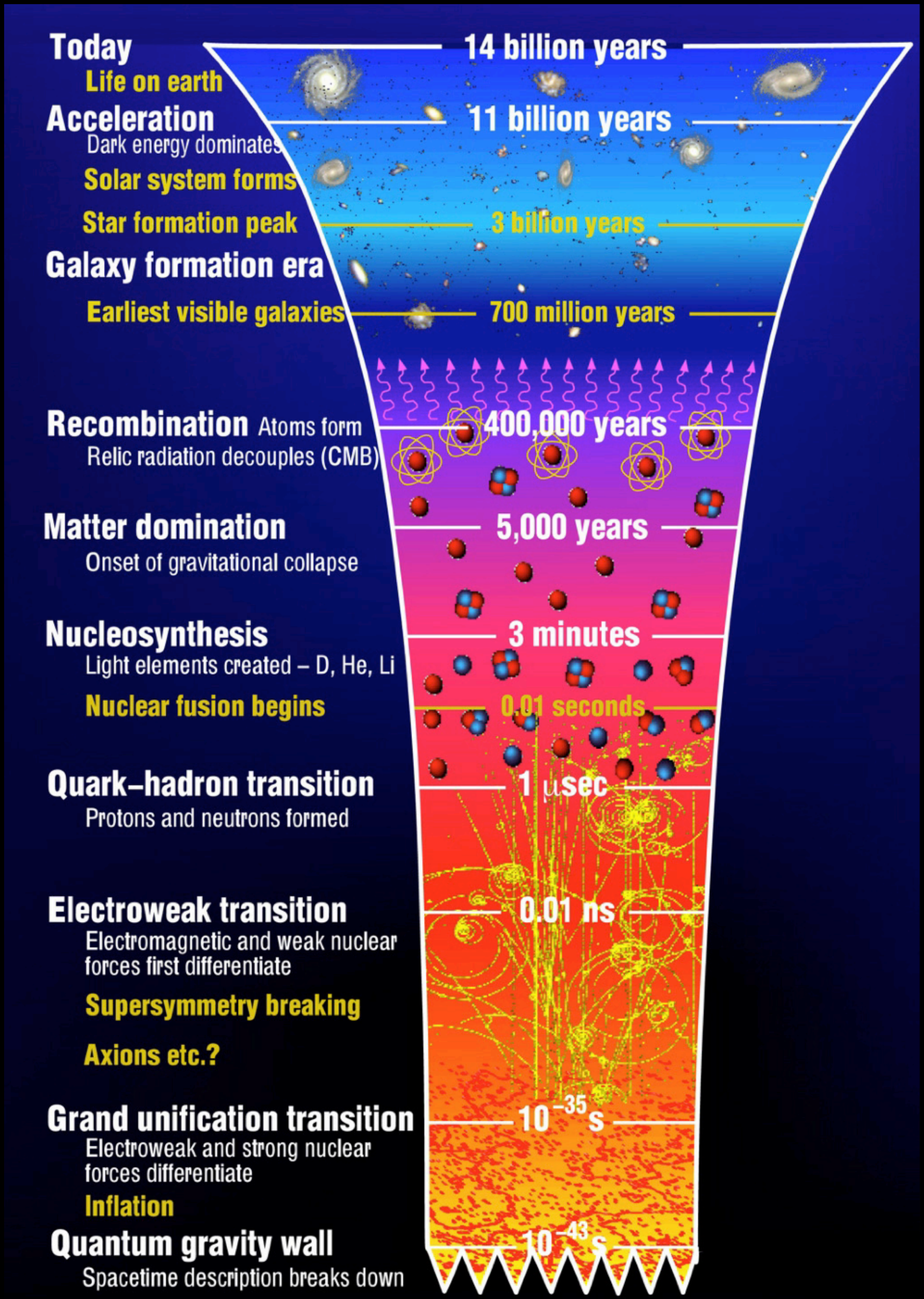
Relic density of dark matter

Using combined data from a variety of telescopes measuring billions of galaxies, CMB and supernovae, we can fit those parameters.

Parameter	TT+lowP+lensing 68% limits	TT,TE,EE+lowP+lensing+ext 68% limits
n_s	0.9677 ± 0.0060	0.9667 ± 0.0040
H_0	67.81 ± 0.92	67.74 ± 0.46
Ω_Λ	0.692 ± 0.012	0.6911 ± 0.0062
Ω_m	0.308 ± 0.012	0.3089 ± 0.0062
$\Omega_b h^2$	0.02226 ± 0.00023	0.02230 ± 0.00014
$\Omega_c h^2$	0.1186 ± 0.0020	0.1188 ± 0.0010
σ_8	0.8149 ± 0.0093	0.8159 ± 0.0086
z_{re}	$8.8^{+1.7}_{-1.4}$	$8.8^{+1.2}_{-1.1}$
Age/Gyr	13.799 ± 0.038	13.799 ± 0.021

Dark matter relic density from cosmological data

What is Relic density?



- Amount of dark matter left over today from the hot dense plasma after Universe expands and cools

- Amount of DM referred to as **relic -** i.e. relic of the early universe

- **Relic density** can also tell us how **DM** was produced in early universe

Why is Relic Density Important?

Any method of DM production in early Universe must match cosmological measurements of relic density

Any theory predicting that dark matter is produced via a certain method, must match relic density from cosmological observations

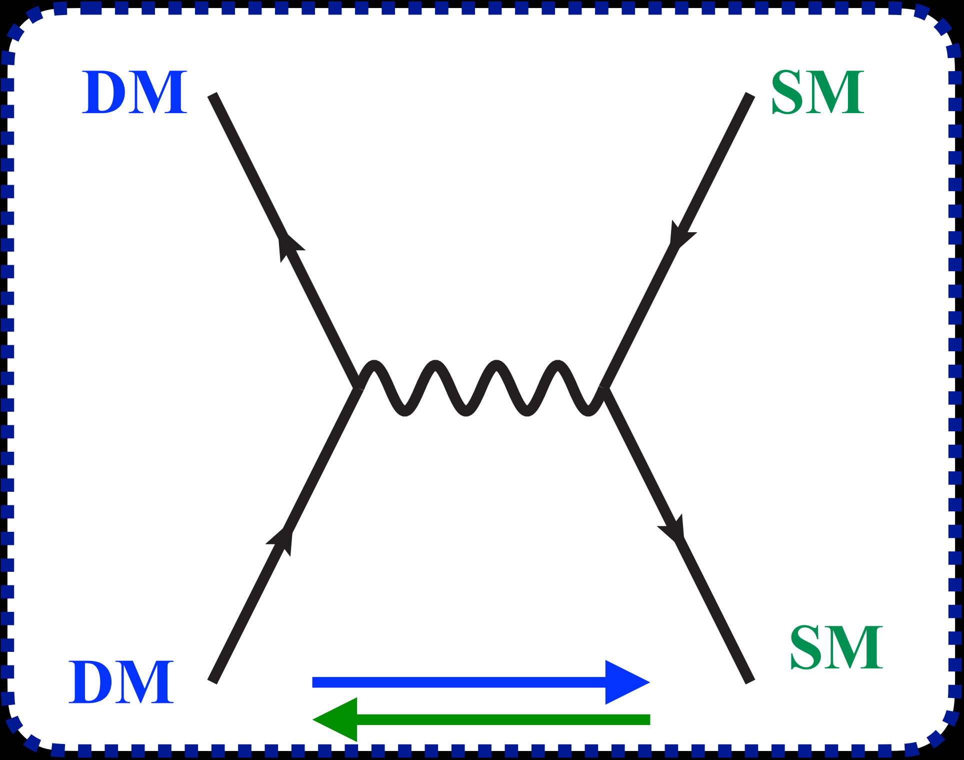
$$\boxed{\Omega_c h^2 \sim 0.12}$$

This is important because it allows the identification of theory parameter space that is interesting and predictive, giving us a clue where to start searching for dark matter

Dark matter production in the early universe

e.g. Thermal Freeze-out

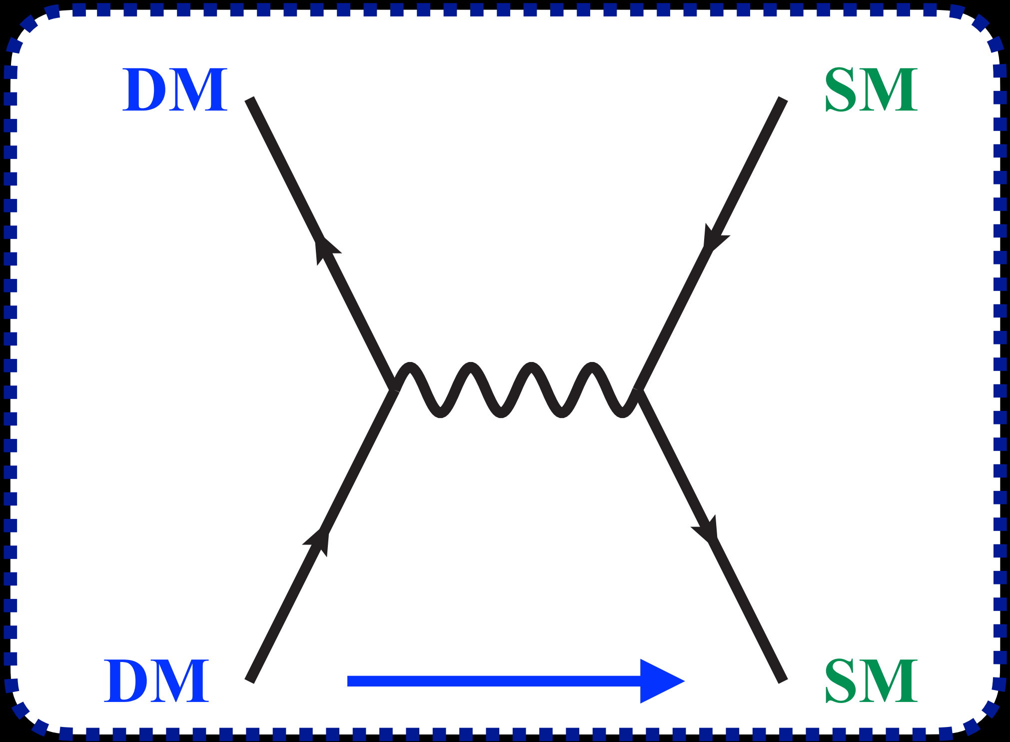
- After big bang, universe is in hot dense plasma of dark sector & SM particles. Plasma is hot enough that



DM and SM are in **thermal equilibrium**

i.e. both particles can be explained by one common temperature

- As universe expands and cools SM cannot convert to DM anymore, only forward process occurs



- As universe keeps expanding, two DM particles cannot find each other to annihilate into SM particle

- **Dark matter has now frozen out and relic number density is set**

Evolution of dark matter number density as Universe expands is given by **Boltzmann equation**

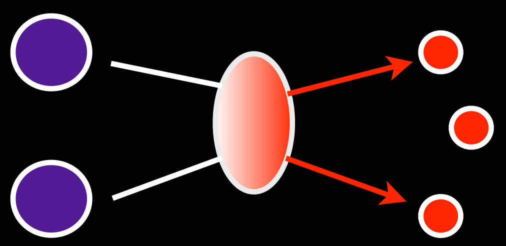
Dark matter number density

$$\frac{dn_\chi}{dt} + 3Hn_\chi = -\langle\sigma v\rangle\left[n_\chi^2 - (n_\chi^{eq})^2\right]$$

Hubble Friction

Annihilation

$H \sim g_* \frac{T^2}{M_{\rm Pl}}$



Describes DM number as universe expands

Evolution of dark matter number density as Universe expands is given by Boltzmann equation

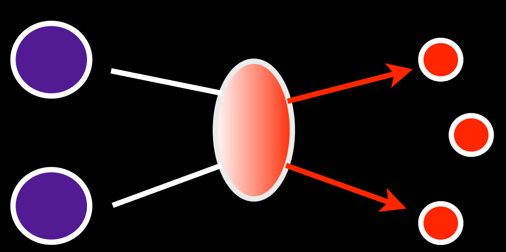
Particles physics enters here

$$\frac{dn_\chi}{dt} + 3Hn_\chi = -\langle \sigma v \rangle \left[n_\chi^2 - (n_\chi^{eq})^2 \right]$$

Hubble Friction

$$H \sim g_* \frac{T^2}{M_{\rm Pl}}$$

Annihilation



Thermally averaged annihilation cross-section: probability that two DM particles traveling in some velocity distribution will find each other and annihilate into SM particles.

Evolution of dark matter number density as Universe expands is given by Boltzmann equation

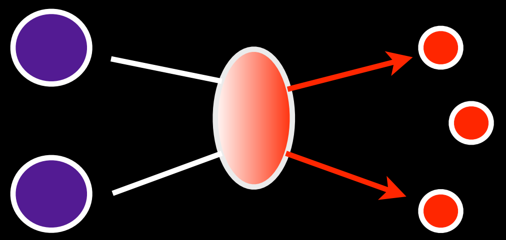
Particles physics enters here

$$\frac{dn_\chi}{dt} + 3Hn_\chi = -\langle\sigma v\rangle\left[n_\chi^2 - (n_\chi^{eq})^2\right]$$

Hubble Friction

$$H \sim g_* \frac{T^2}{M_{\rm Pl}}$$

Annihilation



DM number density at equilibrium, given by

$$n_\chi^{eq} = g\left(\frac{mT}{2\pi}\right)^{3/2} e^{-m/T}$$

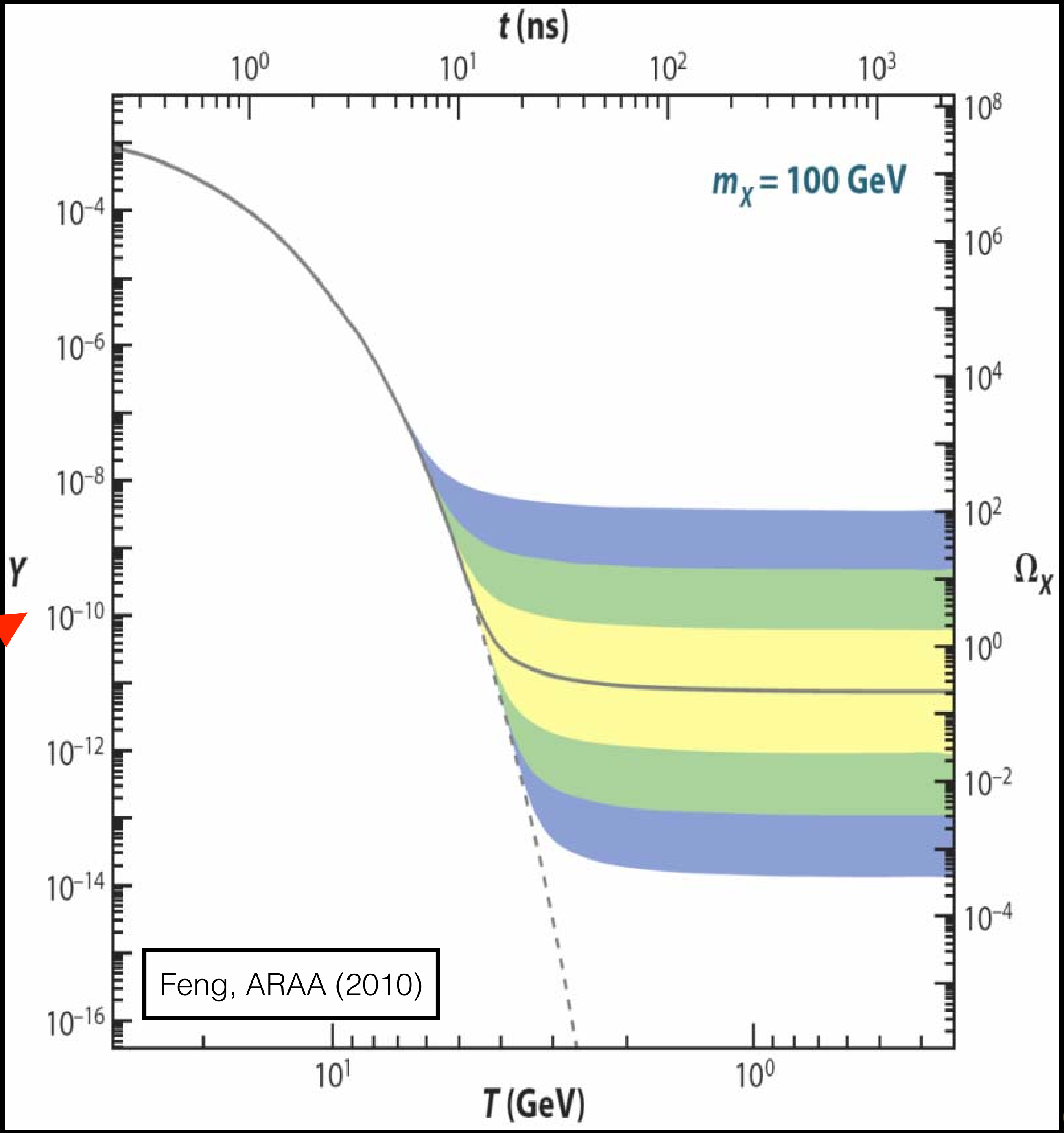
Rewrite **Boltzmann equation** as

$$\frac{dY}{dx} = -\frac{x\,\langle \sigma v \rangle\, s}{H(m)}\left(Y^2 - Y_{eq}^2\right)$$

Comoving number density is $Y = \frac{n_\chi}{s}$

temperature of the universe $x = \frac{m}{T}$

Solution of Boltzmann equation

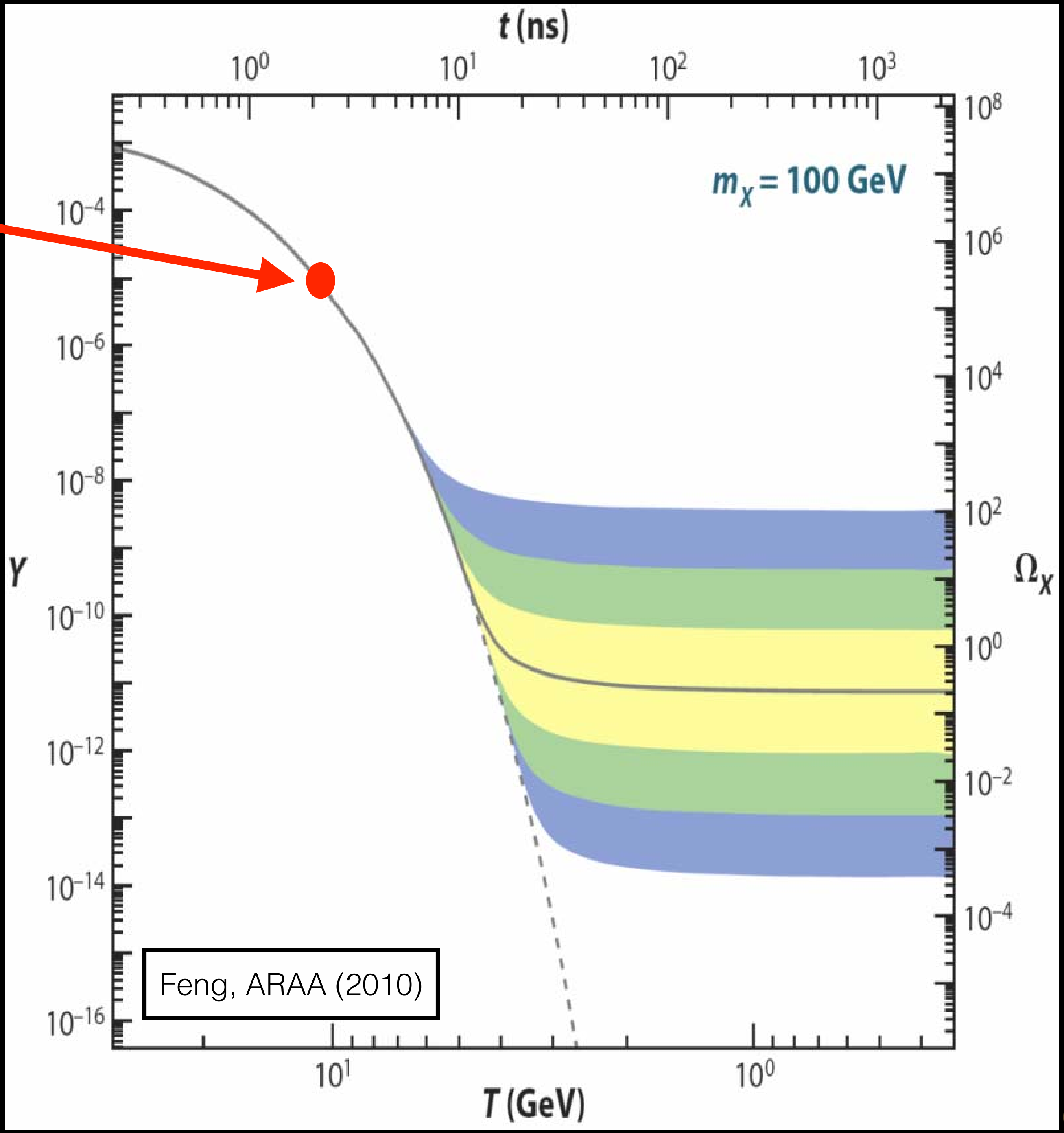


Rewrite Boltzmann equation as

$$\frac{dY}{dx} = -\frac{x\,\langle \sigma v\rangle\, s}{H(m)}\left(Y^2 - Y_{eq}^2\right)$$

DM number density at equilibrium

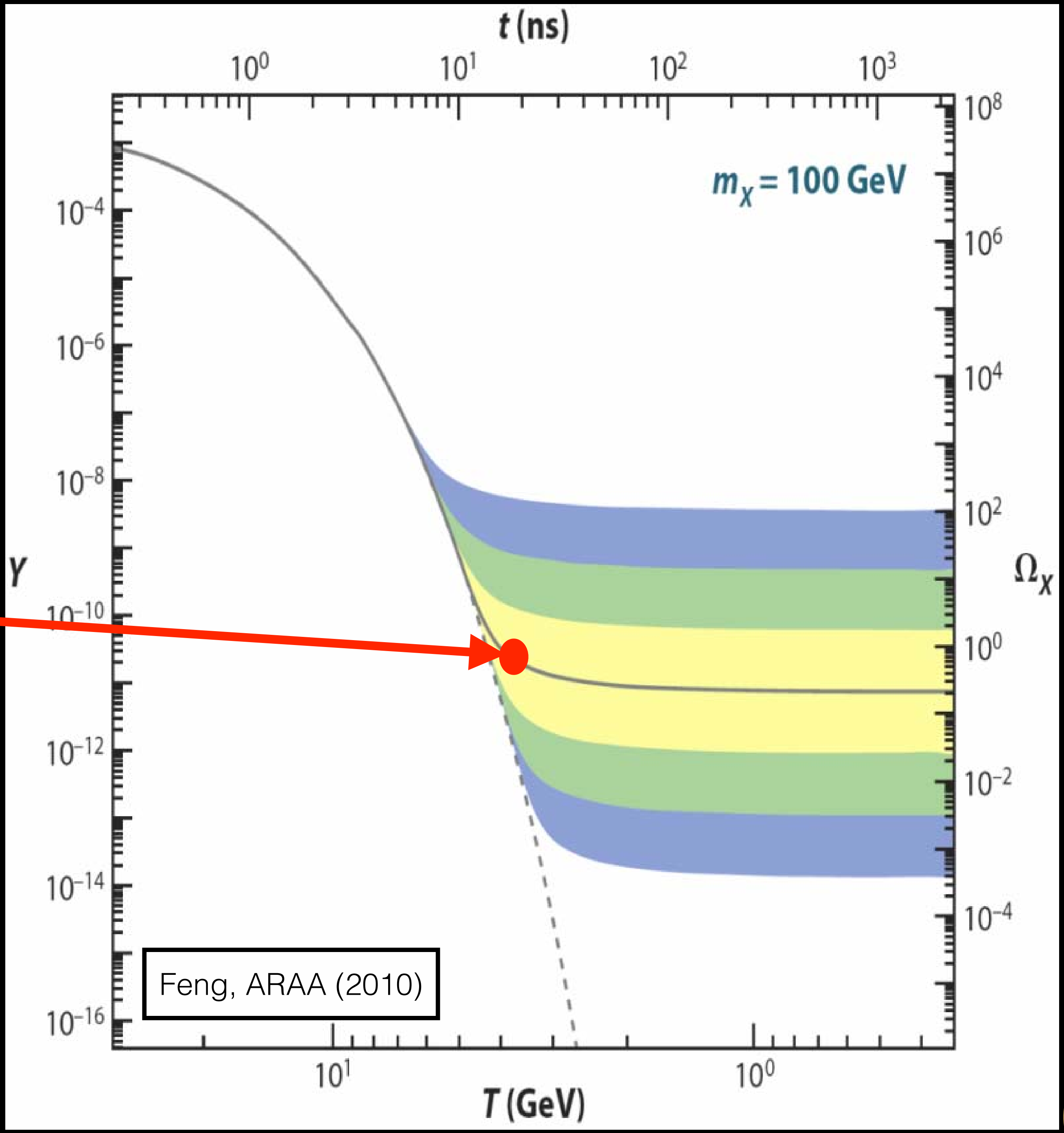
Keeps exponentially decreasing as universe expands and cools and DM converts to SM particles



Rewrite **Boltzmann equation** as

$$\frac{dY}{dx} = -\frac{x\langle\sigma v\rangle s}{H(m)}(Y^2 - Y_{eq}^2)$$

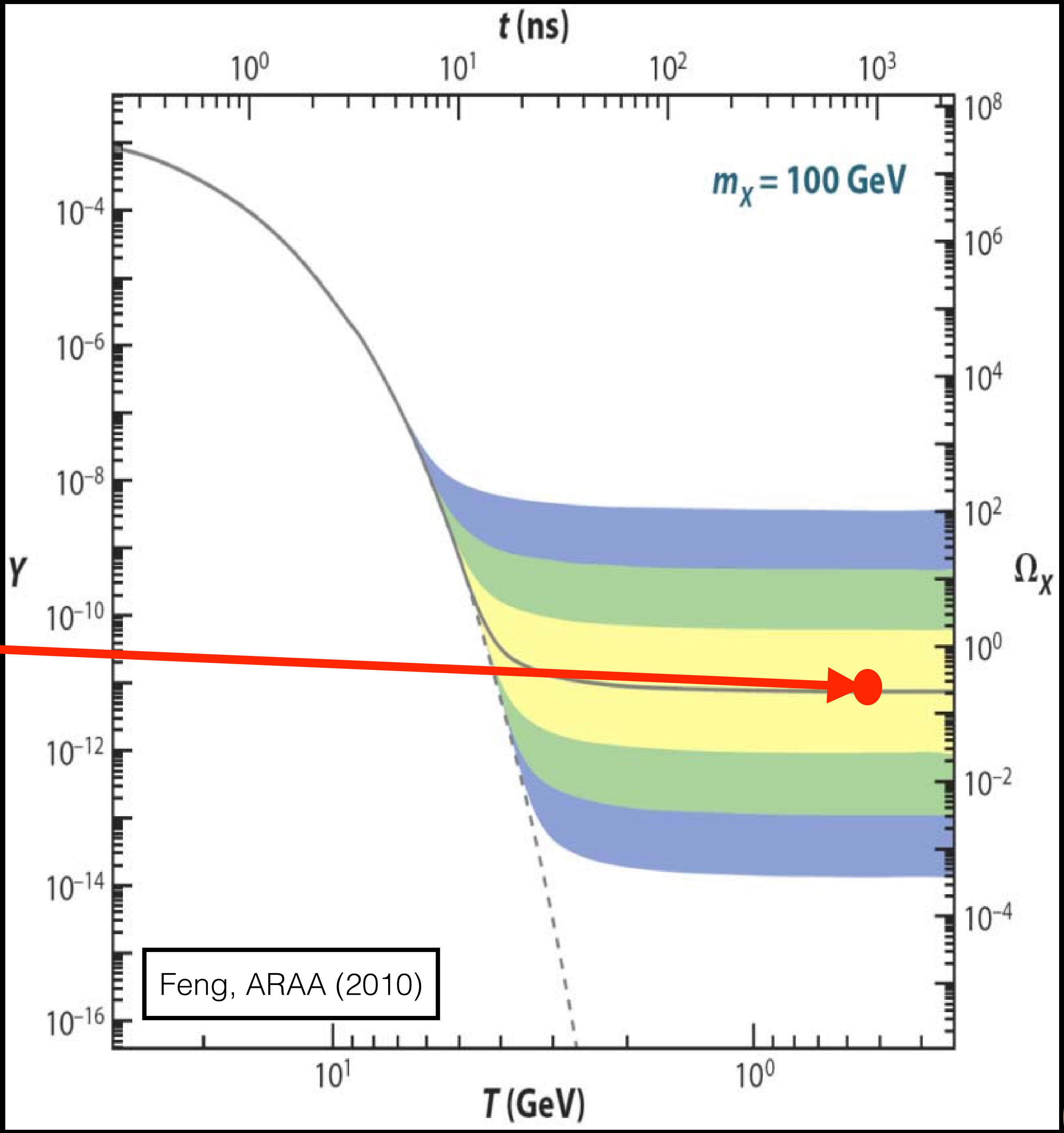
DM Freezes out, particles are no longer able to find each other to annihilate away into SM



Rewrite **Boltzmann equation** as

$$\frac{dY}{dx} = -\frac{x\,\langle\sigma v\rangle\, s}{H(m)}\left(Y^2 - Y_{eq}^2\right)$$

DM number density we observe in the universe today



Rewrite **Boltzmann equation** as

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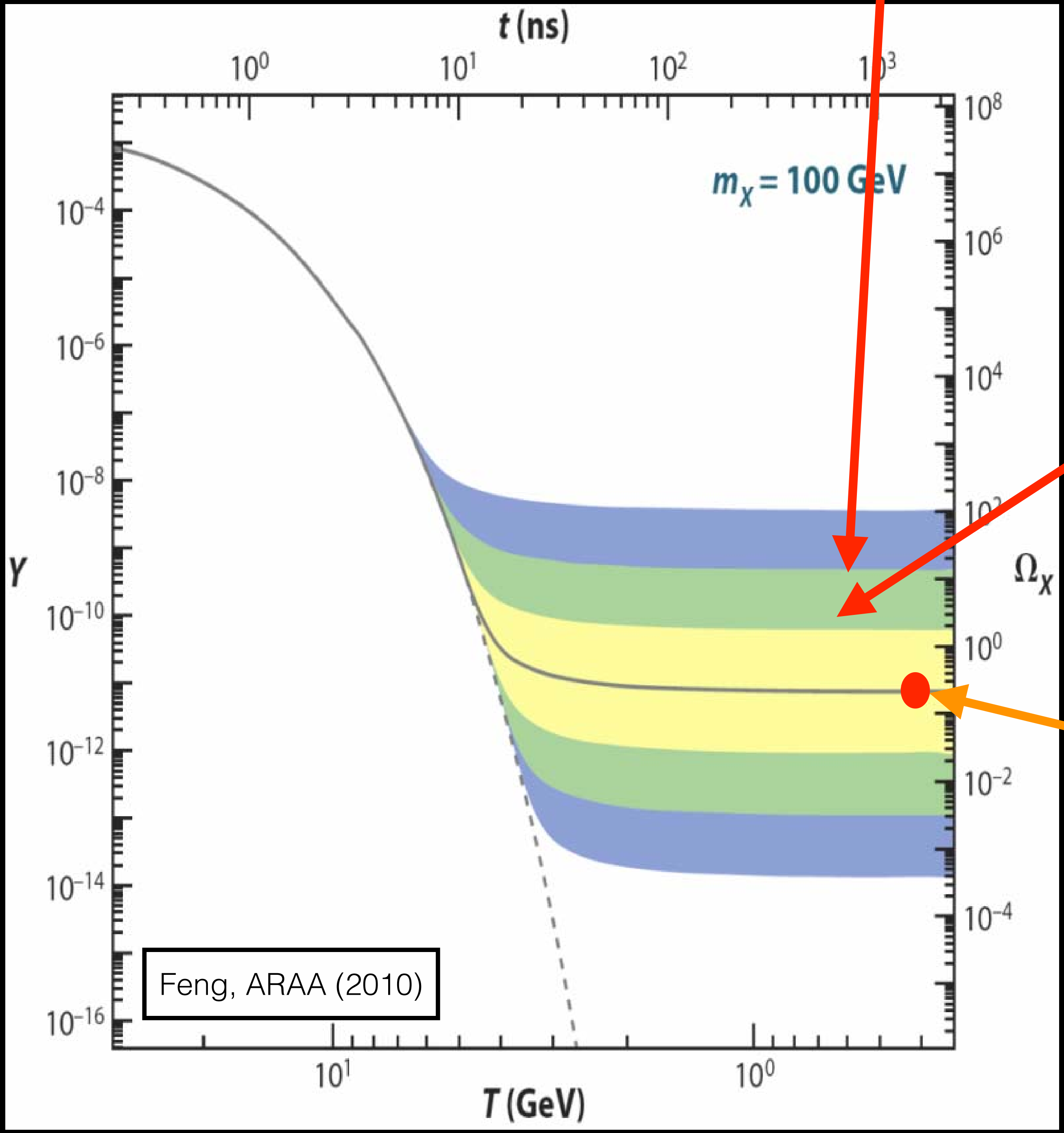


Controlled by strength of particle physics interactions

$\Omega_c h^2 \sim 0.12$

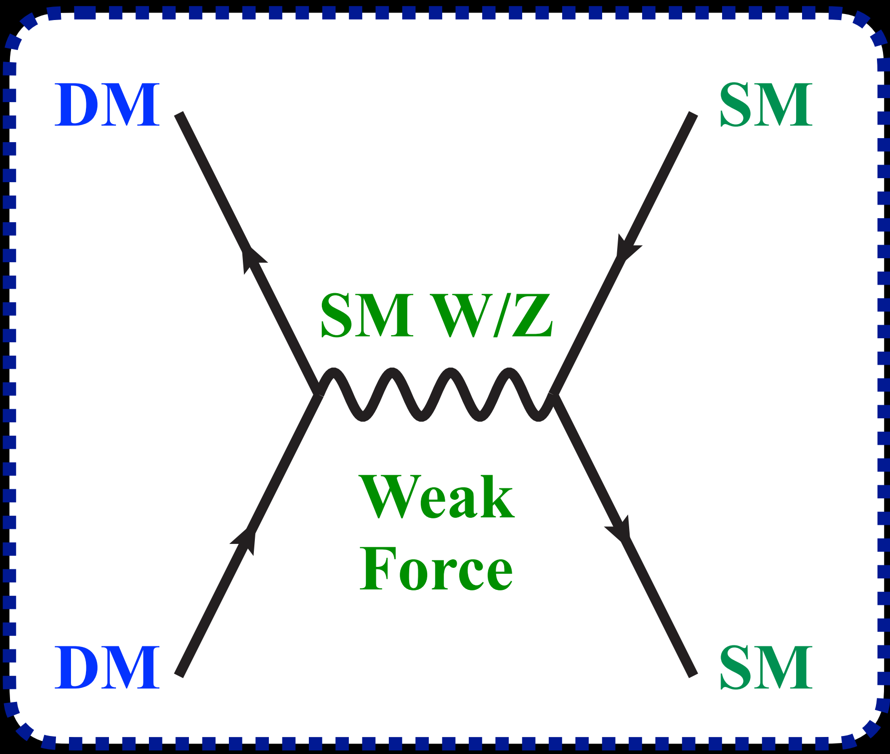
WIMP Miracle

Weakly-Interacting Massive Particles

“Electroweak” interactions ($W^{\pm}$, Z, h)

“weak-scale” mass
1 - 10 000 GeV

Calculations of relic density match cosmological observations almost exactly



Dark matter annihilation

$$\Omega h^2 \sim \frac{0.1\ \mathrm{pb}}{\langle \sigma v \rangle} \sim 0.12$$

Weak scale
annihilation rate

How do we find dark matter?

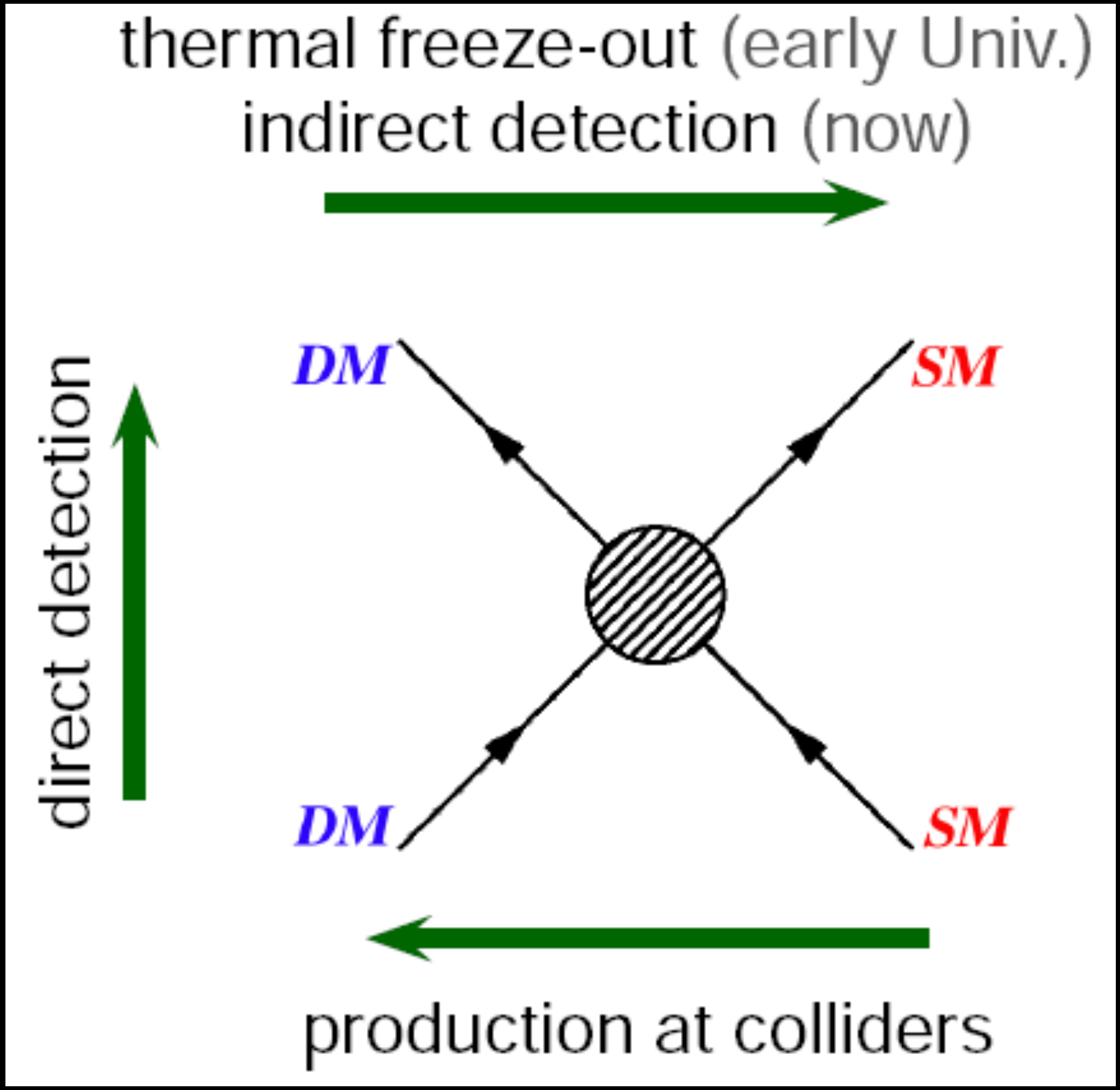


Image: cosmo17.in2p3.fr

In the Sky

At Colliders

In underground detectors

Indirect detection

In Dark matter dense regions like the Galactic center:

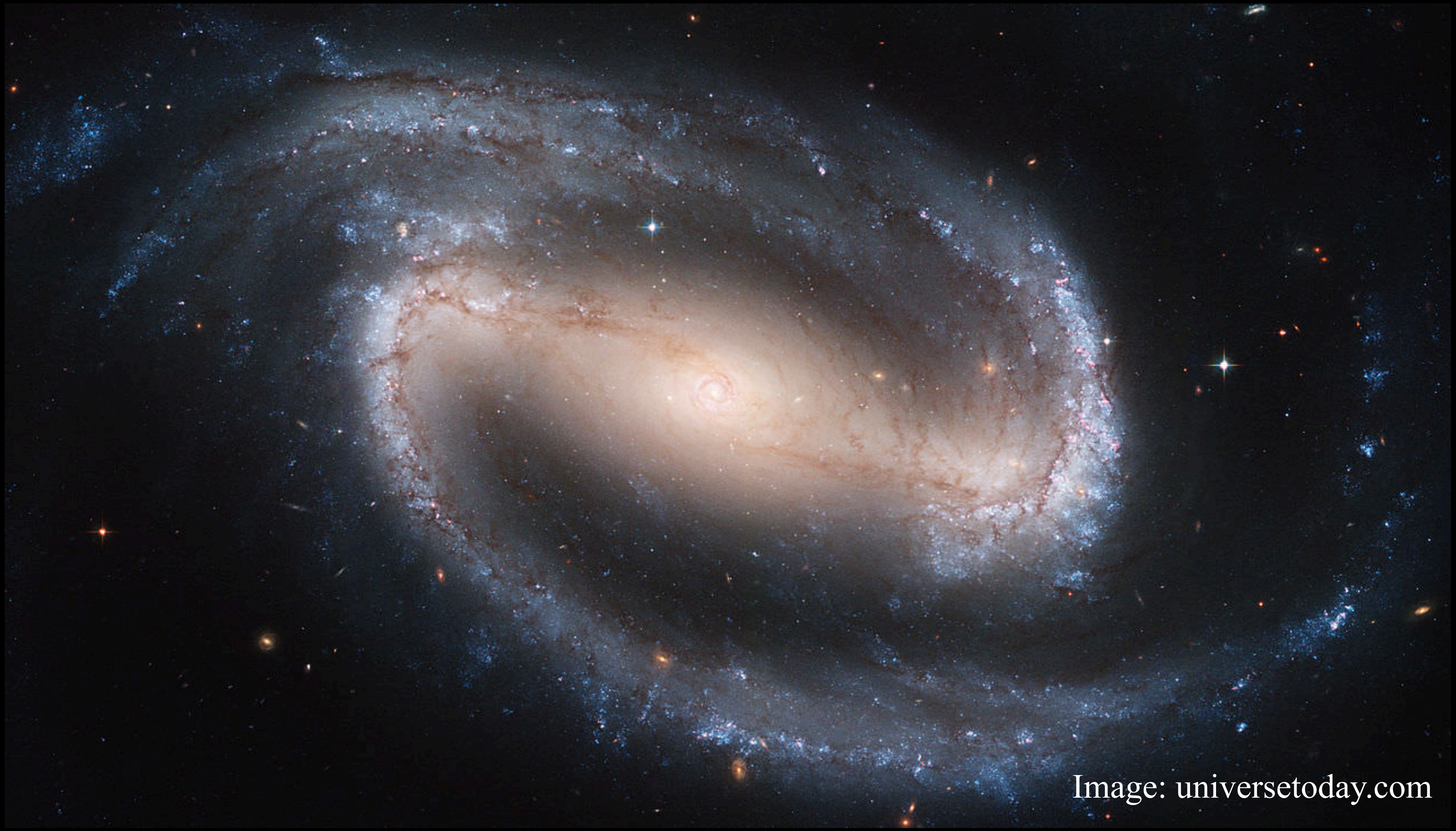


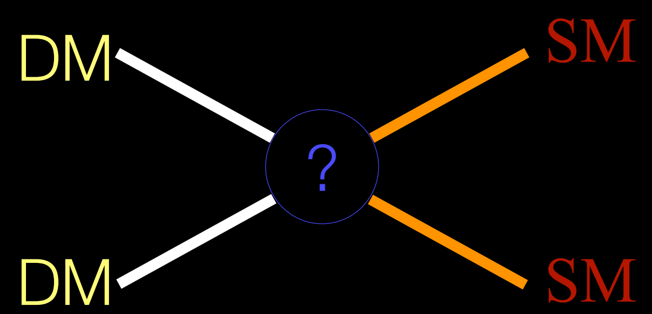
Image: universetoday.com

Dark matter particles find each other and annihilate into SM particles

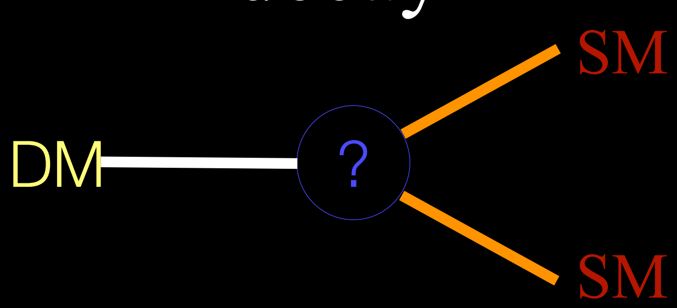
Dark matter particles may decay into SM particles

We must look where dark matter density is highest

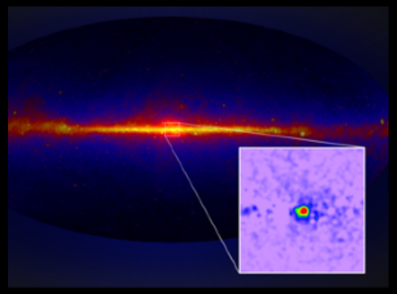
annihilation



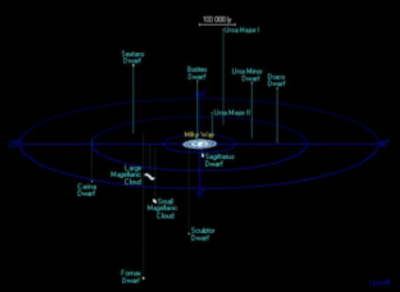
decay



Local halo



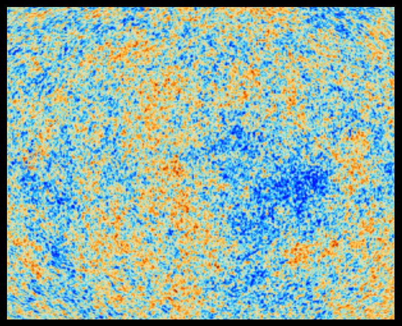
Galactic centre, other galaxies



Dwarf galaxies

The Sun

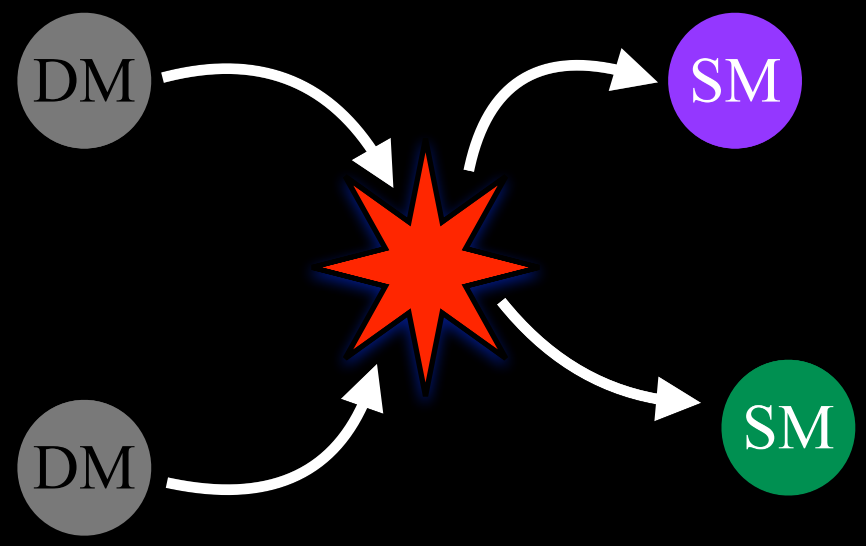
Galaxy Clusters

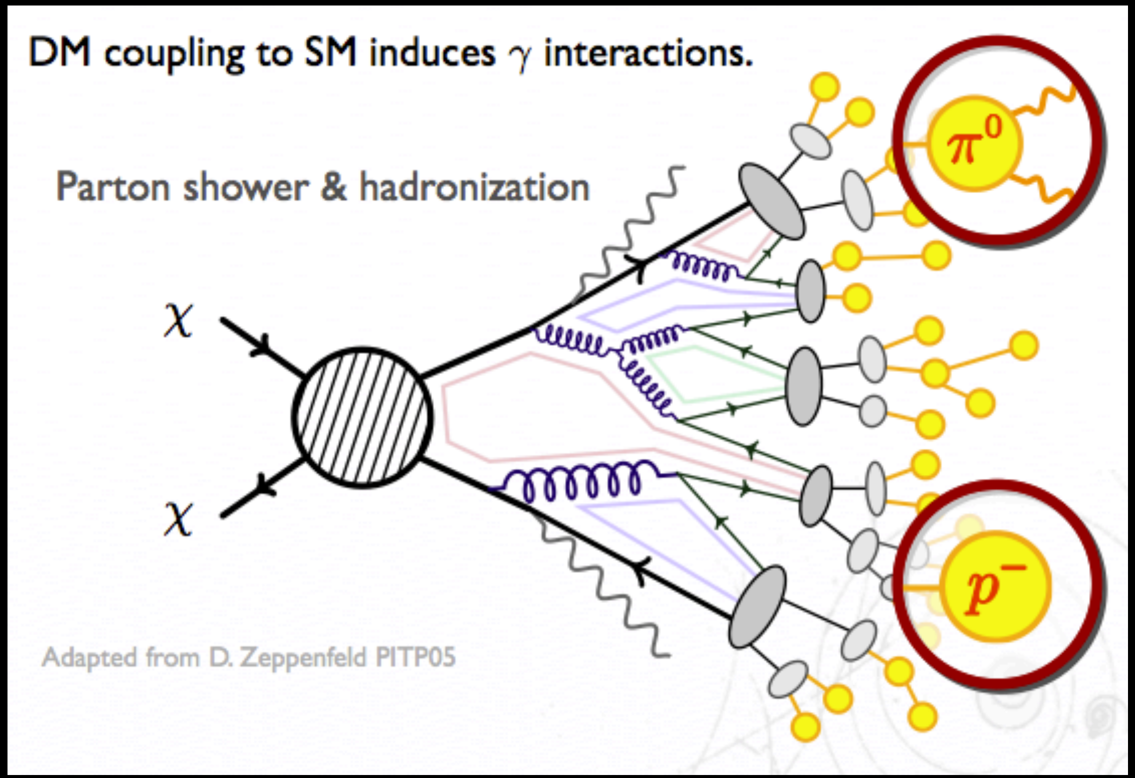


The entire Universe

Indirect Detection

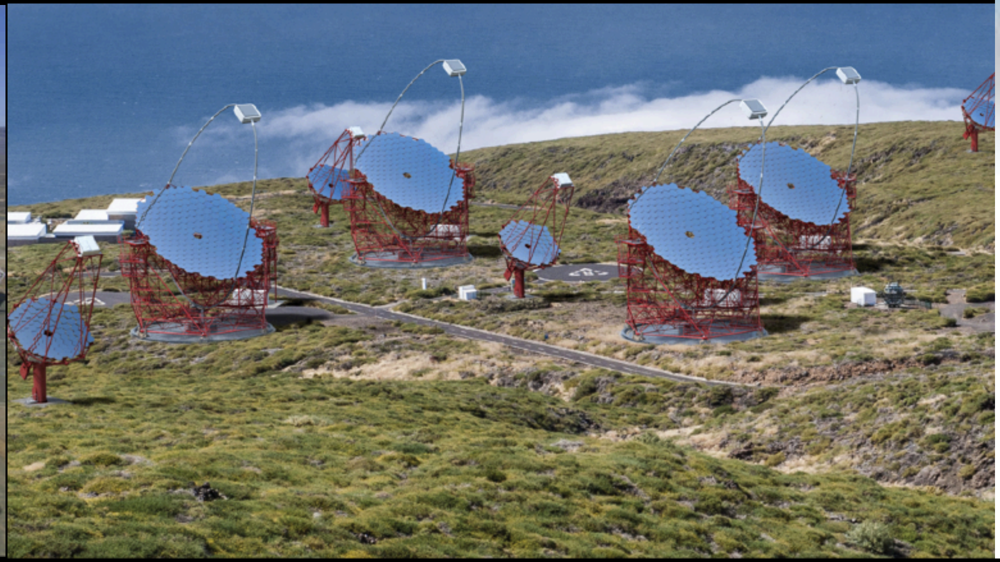
Dark Matter annihilates / decays into SM particles



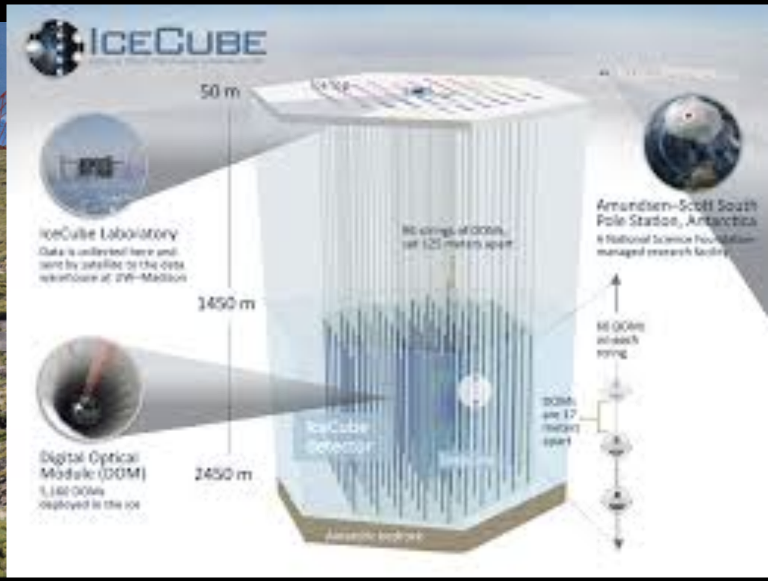


Square Kilometer Array

H.E.S.S / Cherenkov Telescope Array

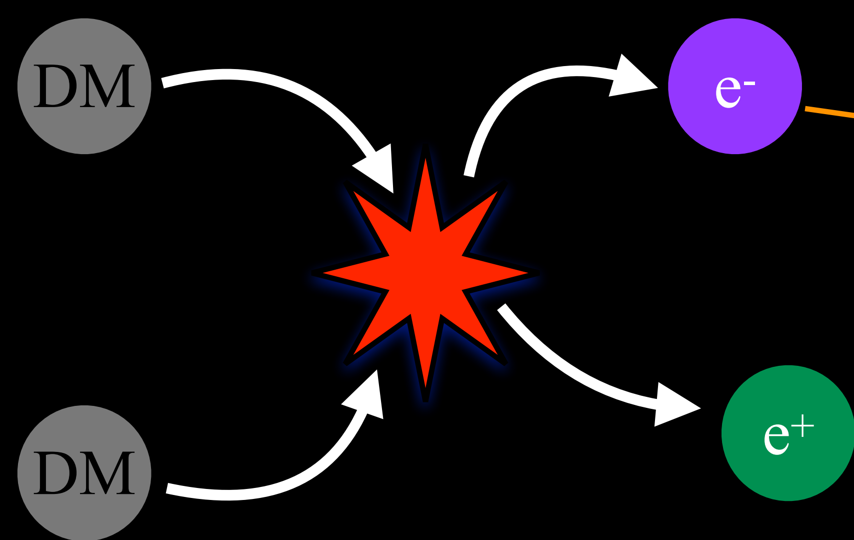


IceCUBE neutrino observatory

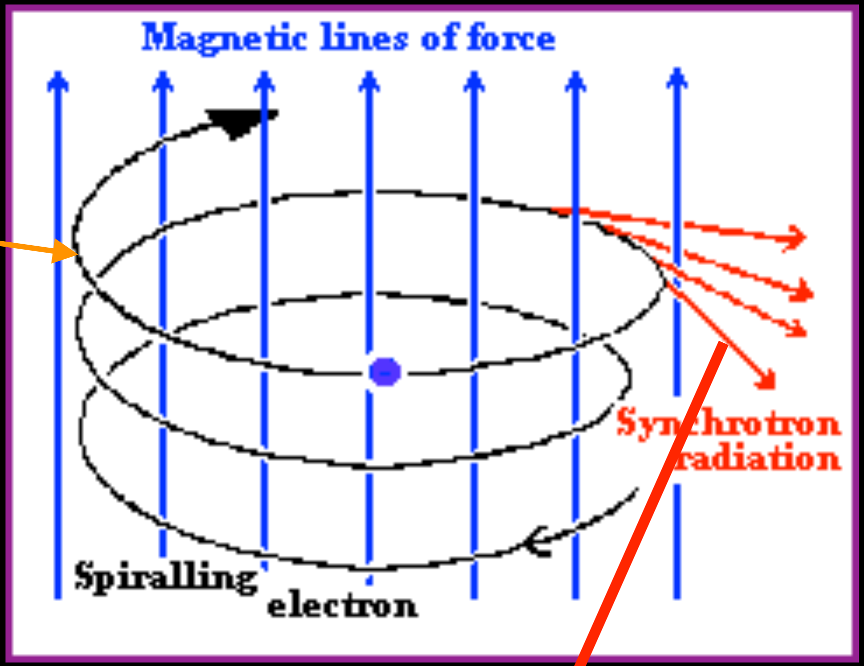


e.g Synchrotron Radiation

Dark Matter annihilates to charged SM particles like electrons/ positrons



Galactic magnetic fields



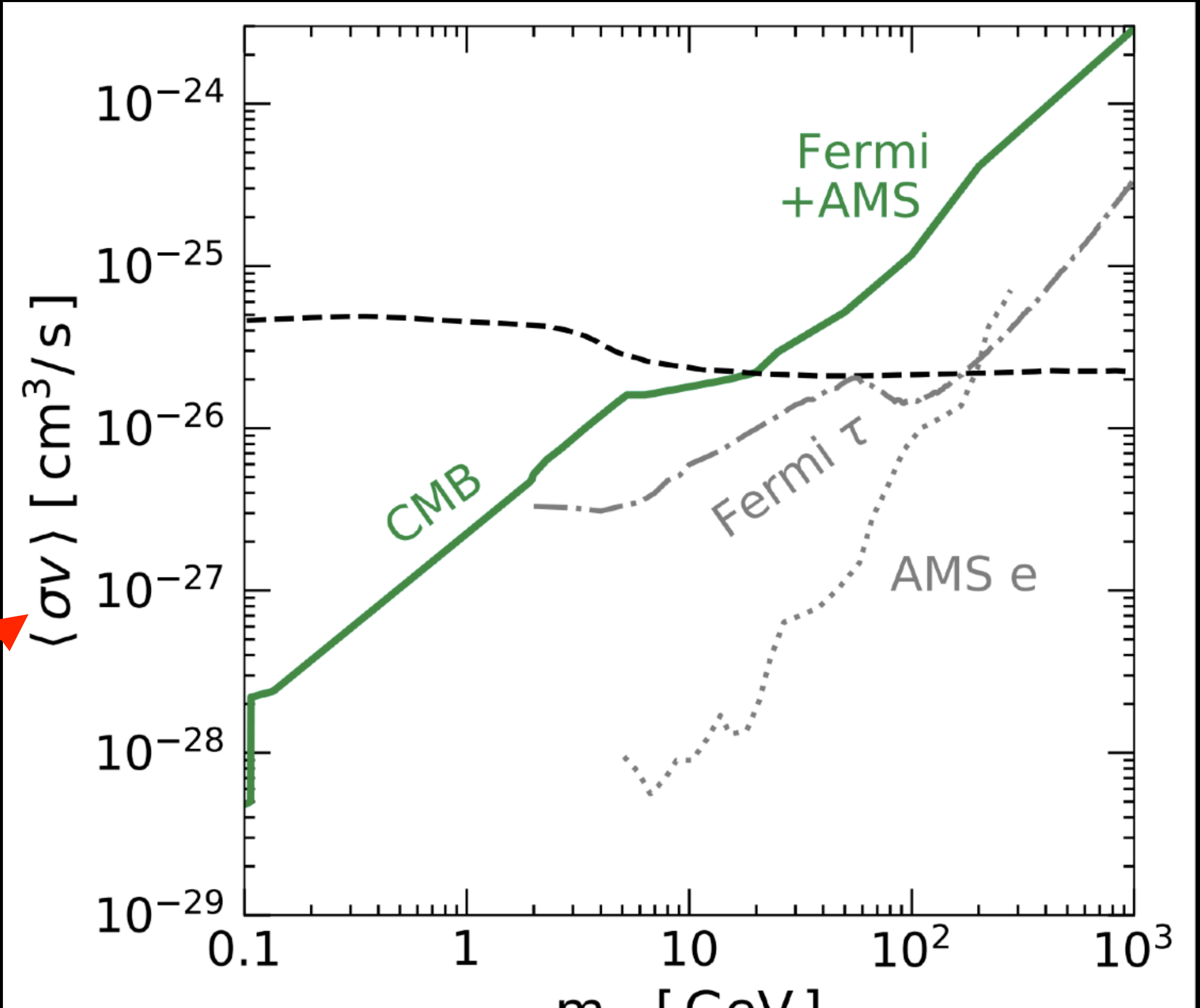
Square Kilometer Array

Indirect Detection

$$\frac{d\phi}{dE_\gamma} \sim \left(\frac{\langle \sigma v \rangle}{8\pi} \frac{dN_\gamma}{dE_\gamma} \frac{1}{m_\chi^2} \right) \int_{\Delta\Omega} \int_{l.o.s} \rho_\chi^2(l) dl d\Omega$$

Flux of photons from DM annihilation

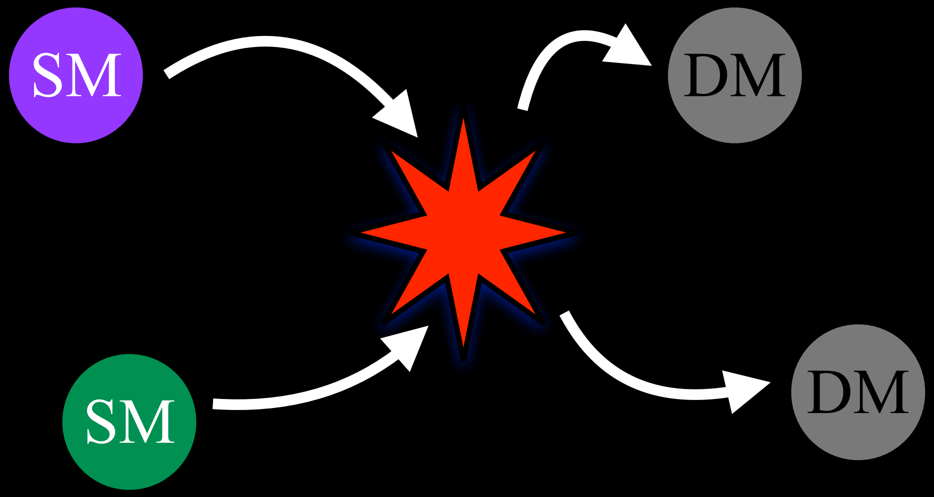
Compared with data from telescopes to obtain limits on DM parameters

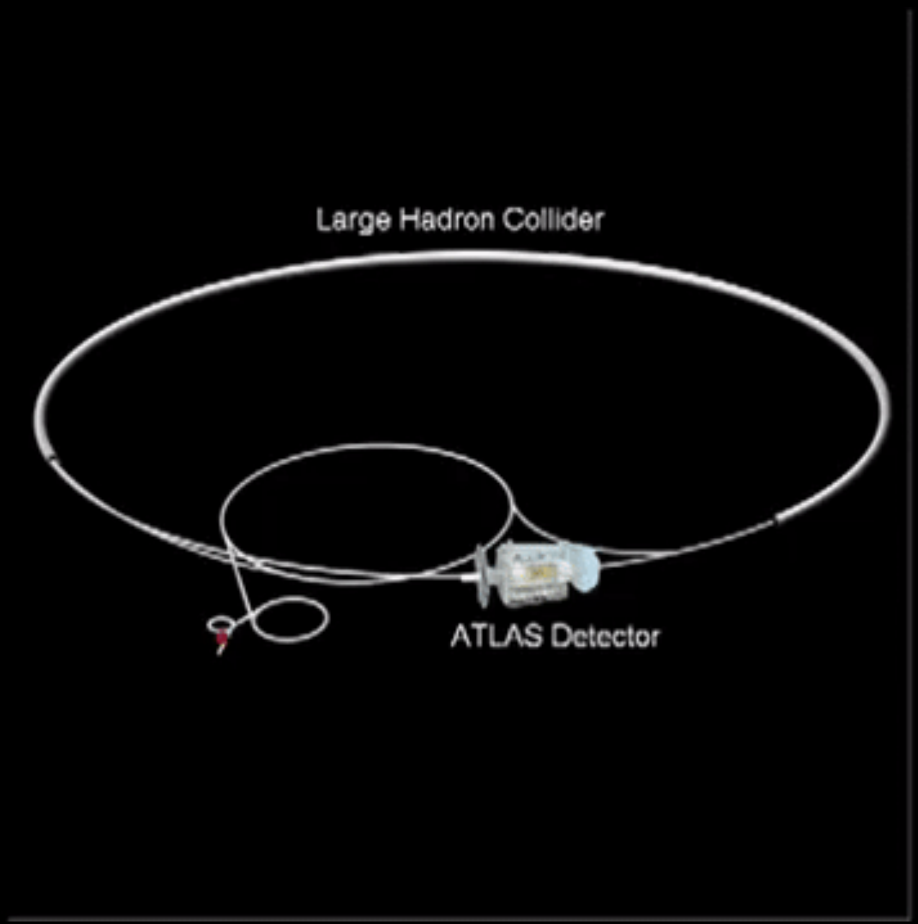


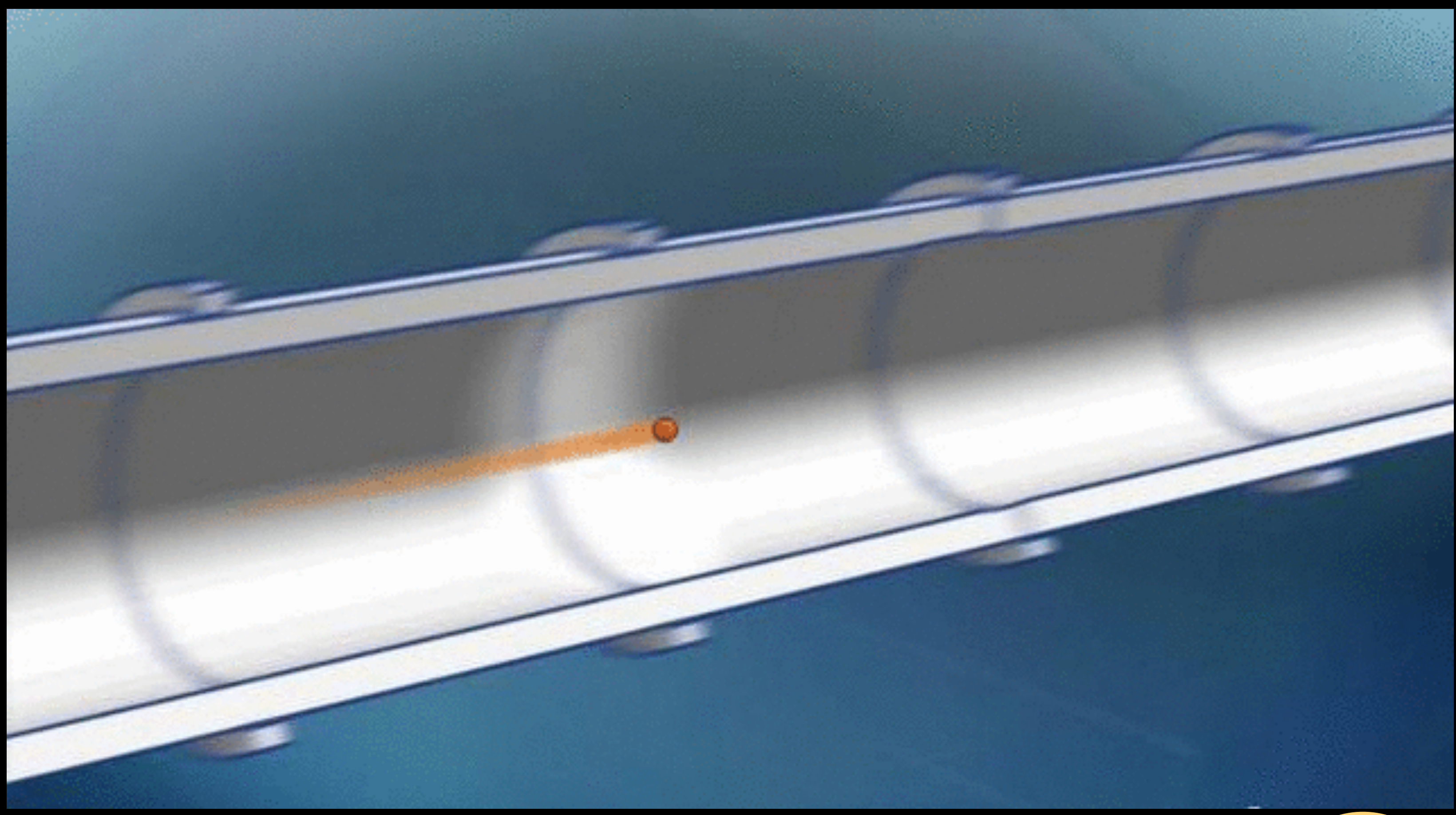
Leane et al, Phys.Rev.D 98 (2018) 2

Production at Colliders/Accelerators

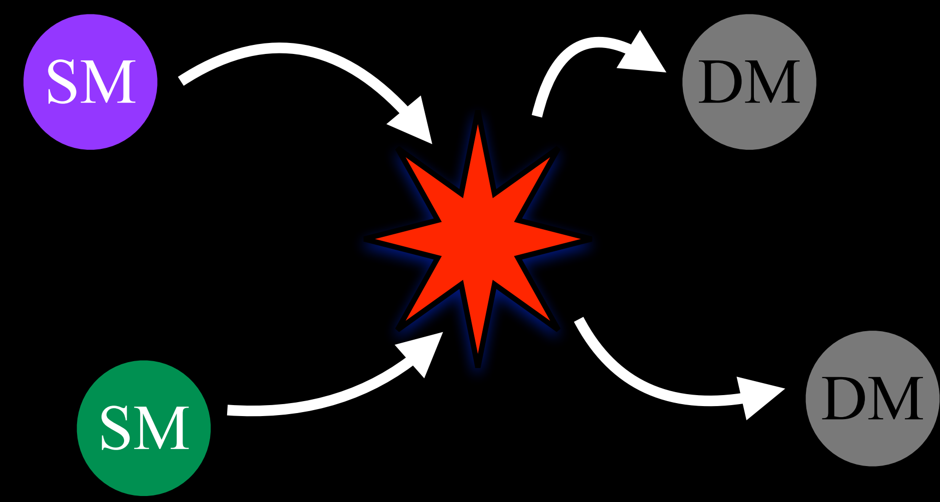
Collide SM particles to produce dark matter in the Laboratory







Production at Colliders/Accelerators



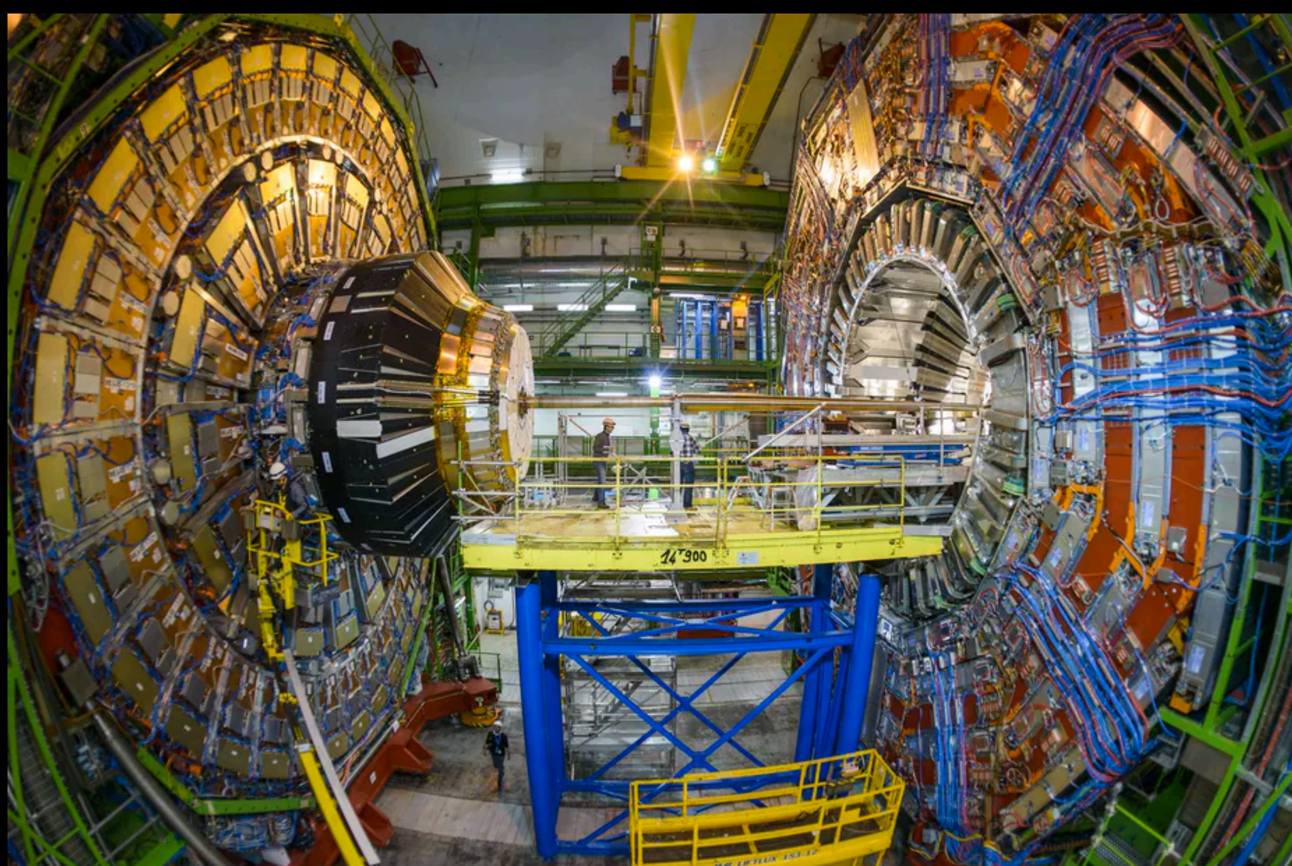


Image: theconversation.com

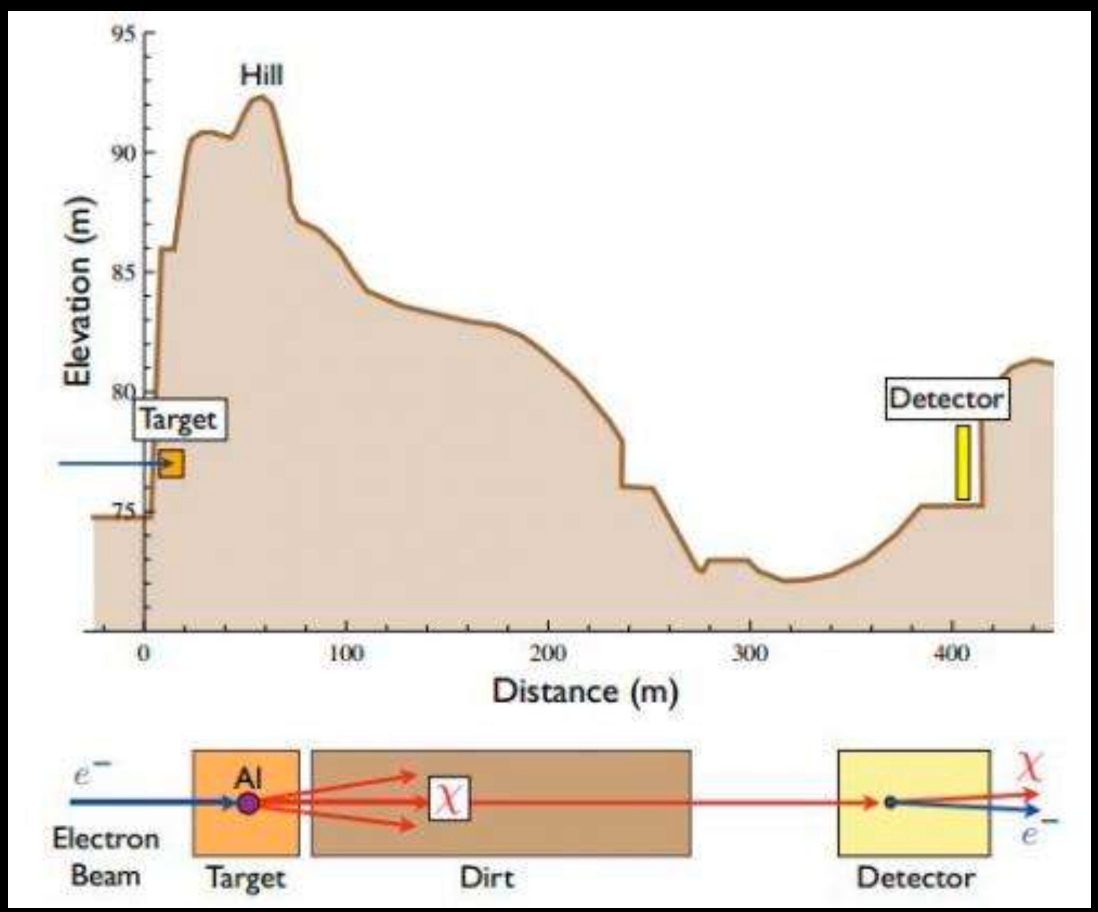


Image: phys.org

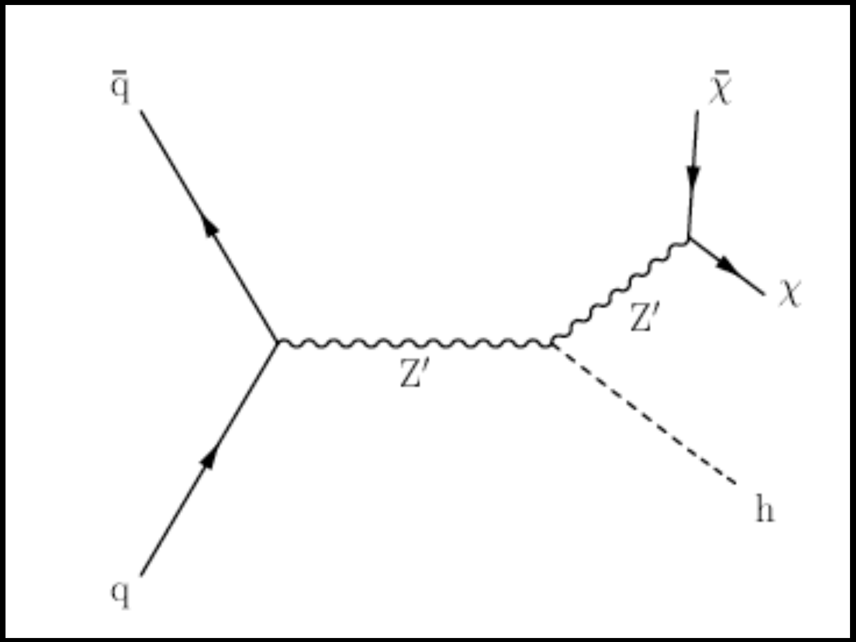
High/low energy colliders

Fixed target experiments

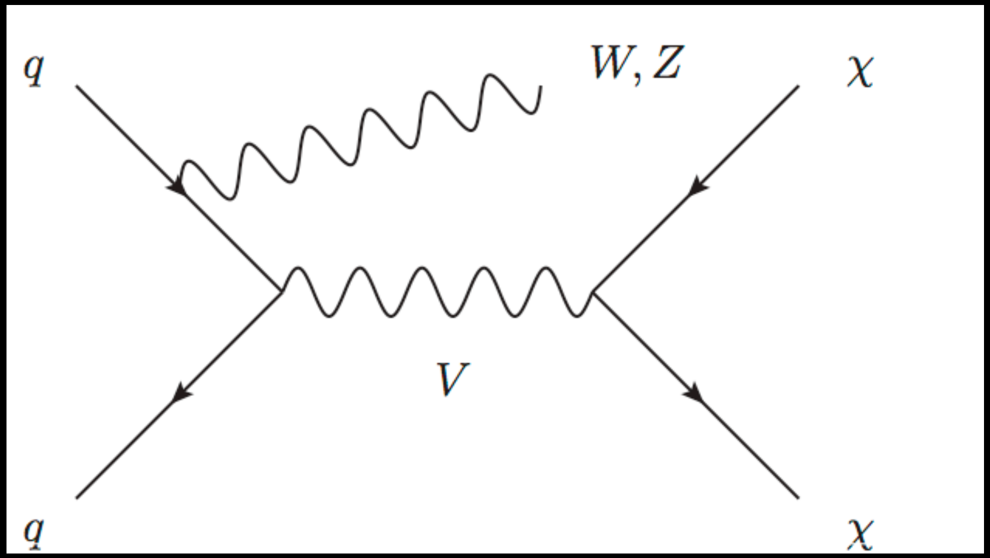
Production/Detection at Colliders

For center of mass collisions, dark matter can be discovered through Mono-X searches

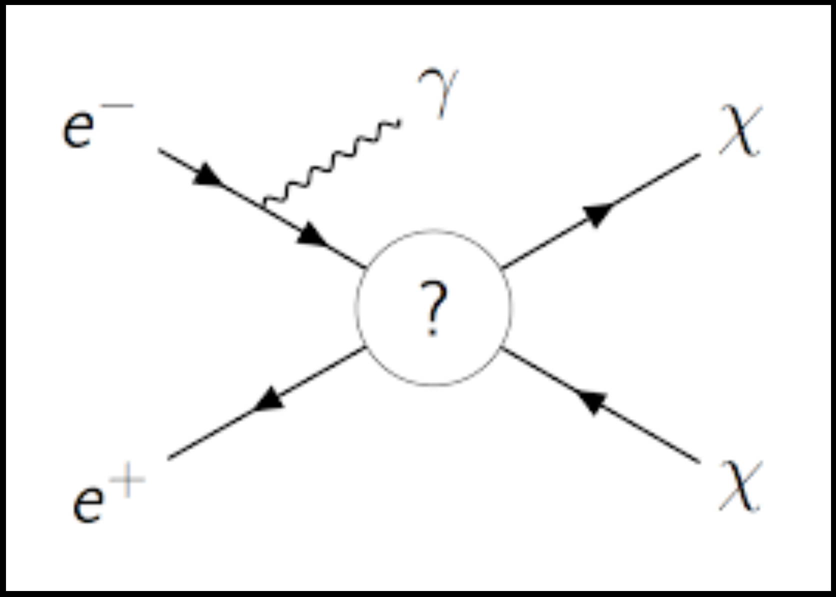
DM is produced and recoils against SM particle, we only see the SM particle. Kinematics allows us to determine DM mass and coupling strength



E.g. mono-Higgs searches at P-P colliders like LHC



mono-W/Z searches at P-P colliders like LHC



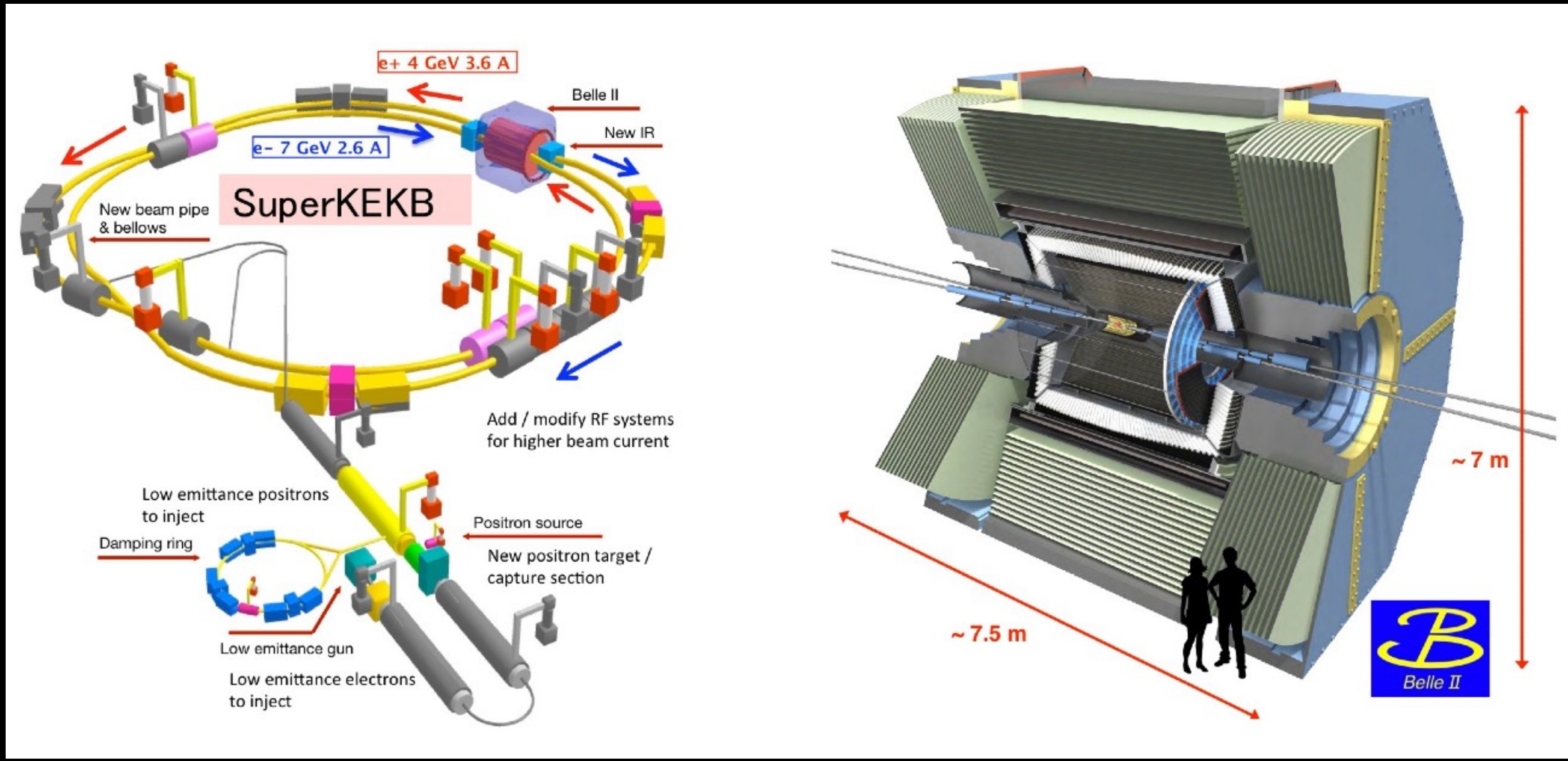
mono-photon searches at e+/e- colliders

Large Hadron Collider at CERN

14 TeV P-P collisions

Higher DM masses

Belle II Collider at KEK in Japan

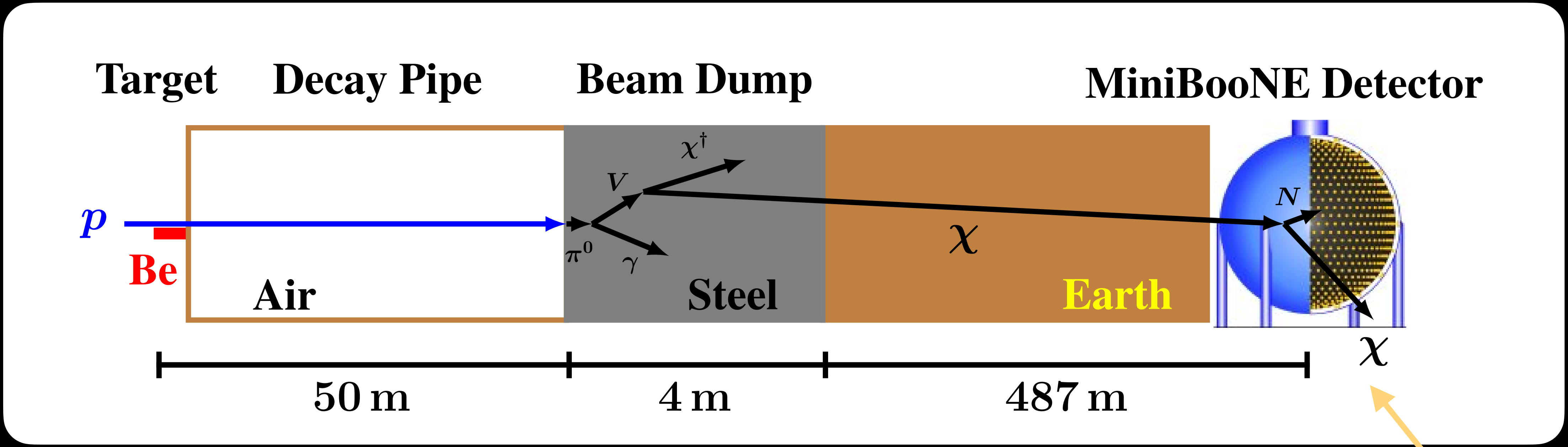


10 GeV e^+/e^- collisions

Probe low DM masses

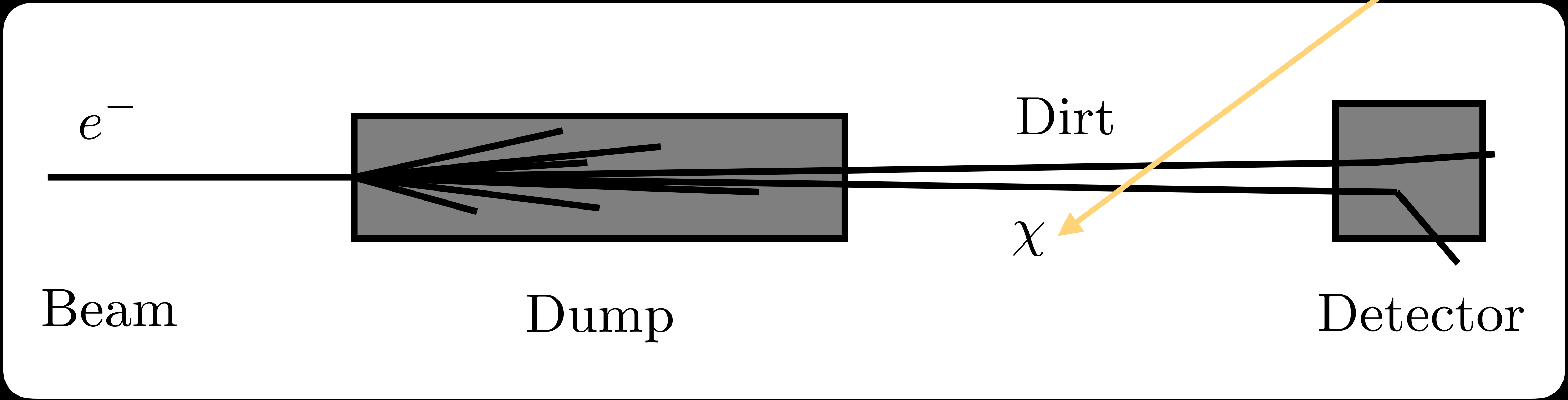
Production/Detection at Fixed target experiments

Proton fixed target experiments



Dark Matter

Electron fixed target experiments



Search for low mass dark matter

Direct Detection

Milky way is surrounded by 'spherical' halo of DM

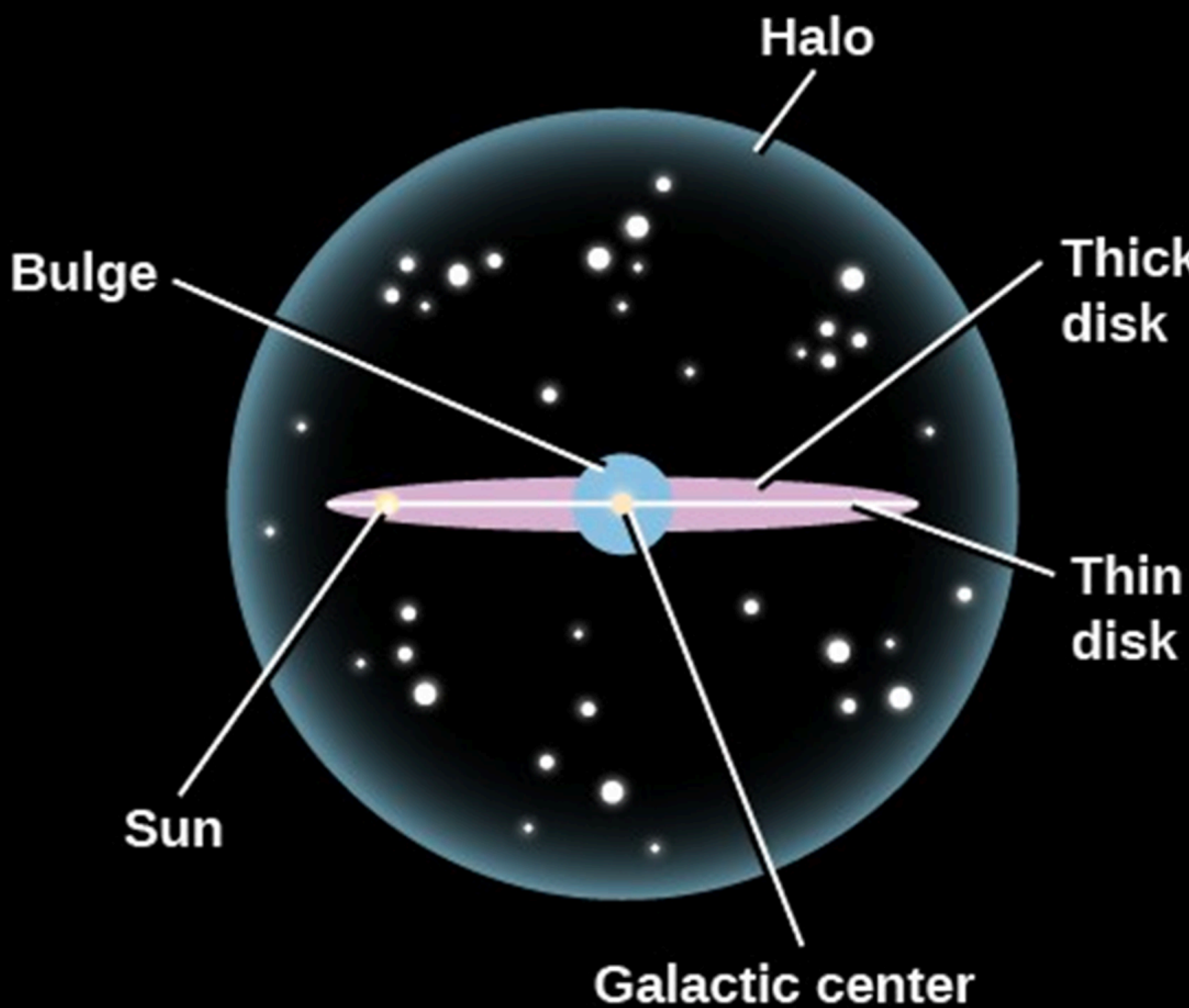


Image: evolution.calpoly.edu

Direct Detection

As sun moves around galaxy, solar system gets hit by dark matter wind

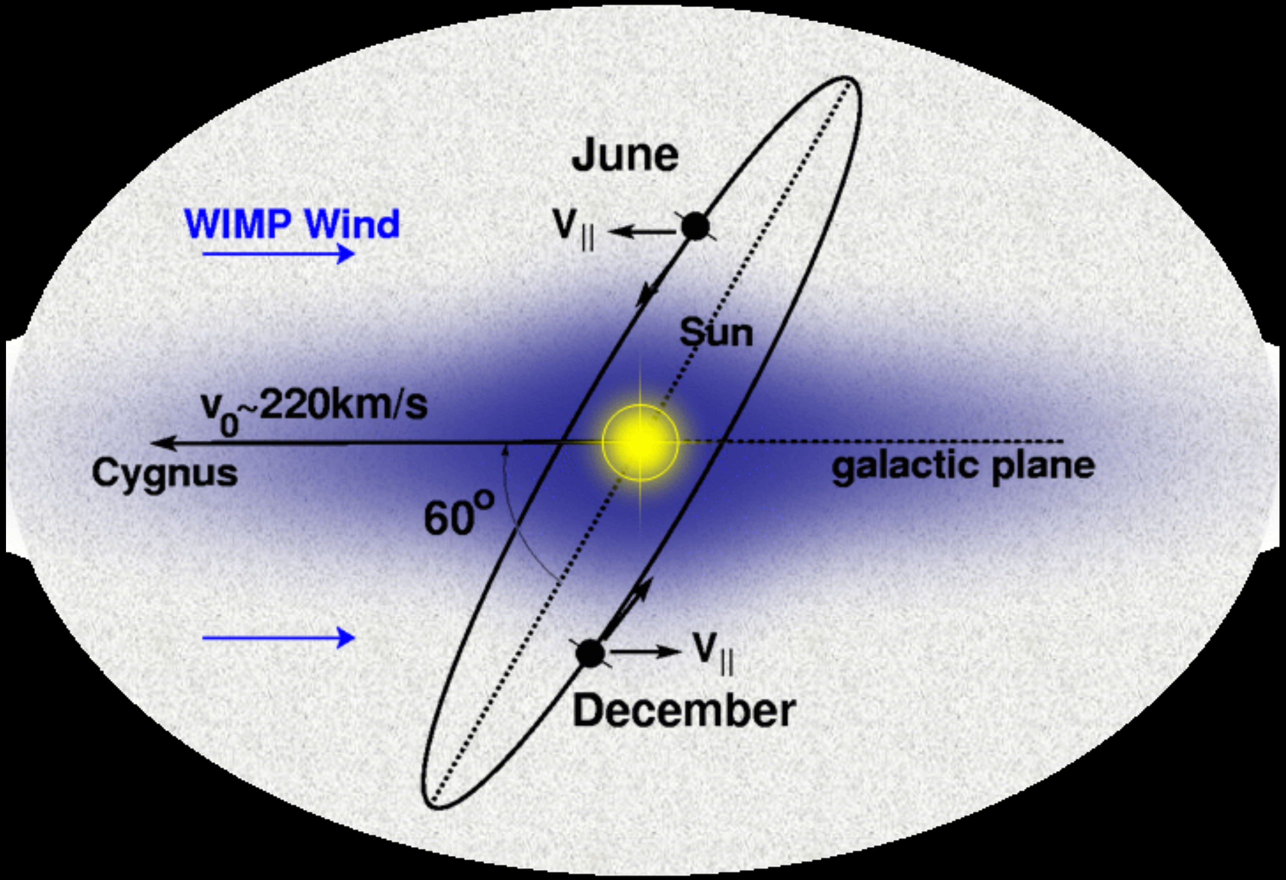


Image: quantumdiaries.org

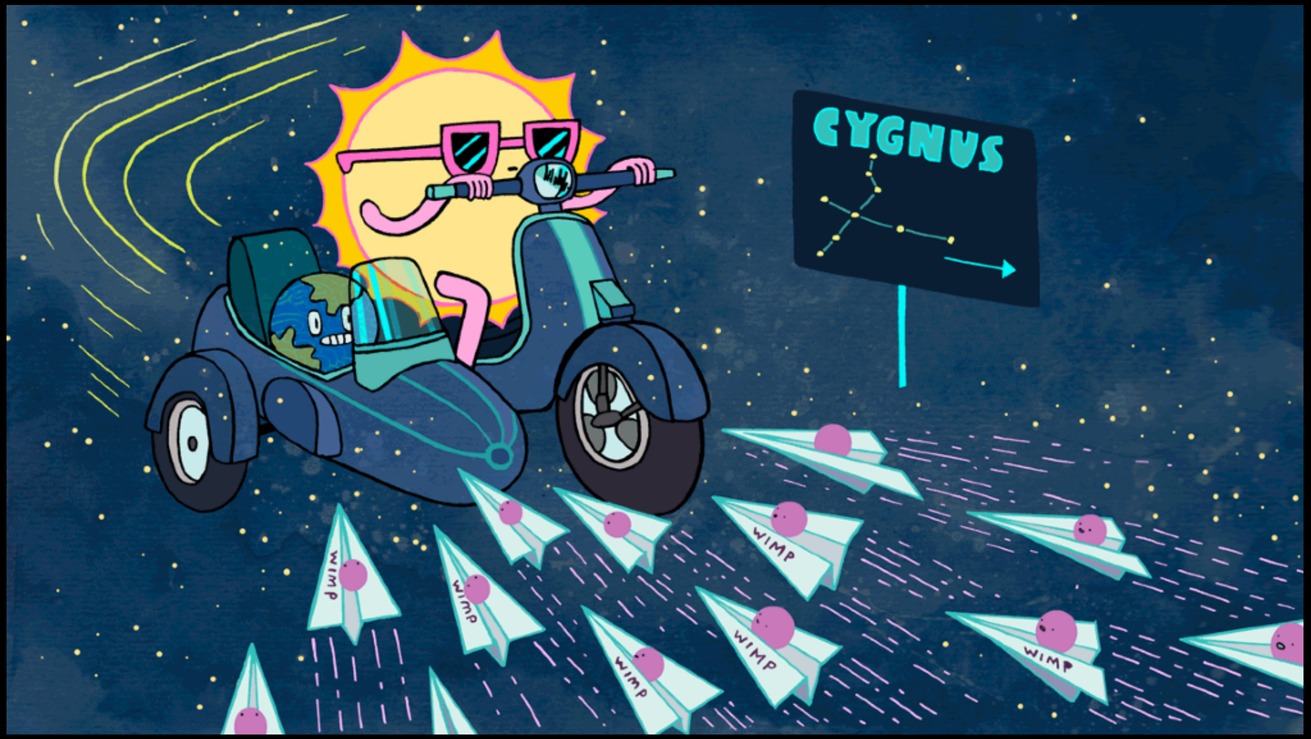


Image: symmetrymagazine.org

Build a detector in a quiet place and patiently wait for dark matter to come knocking

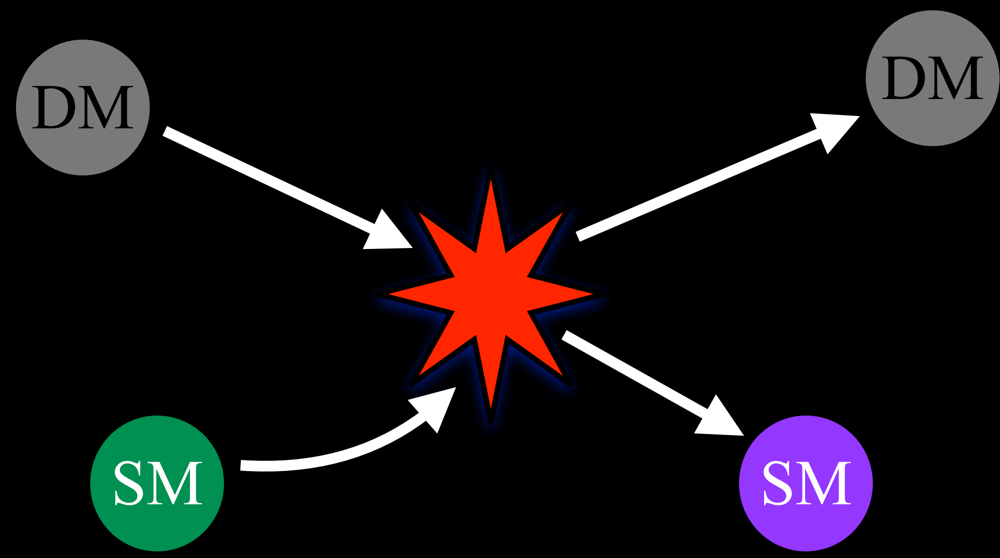
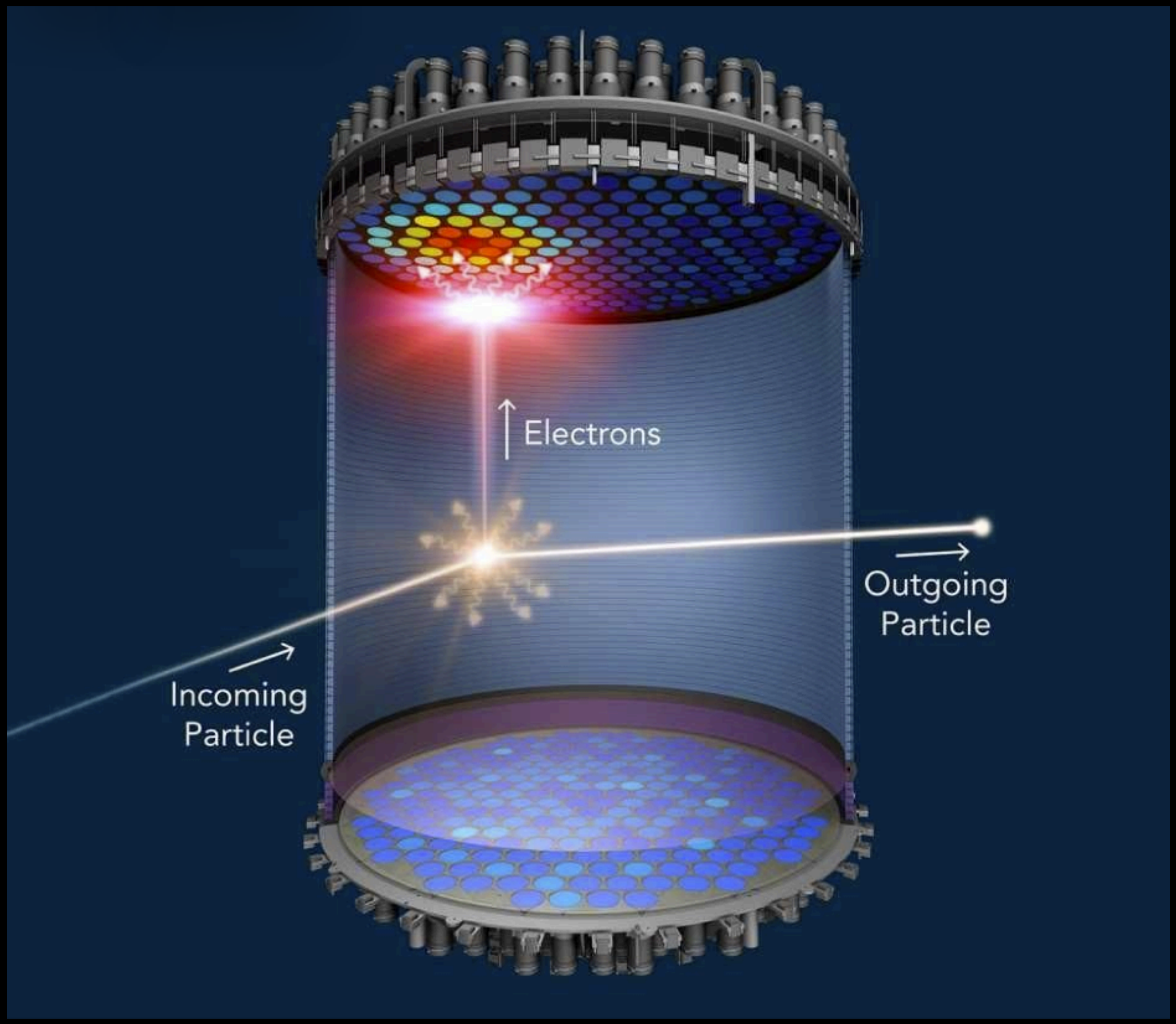


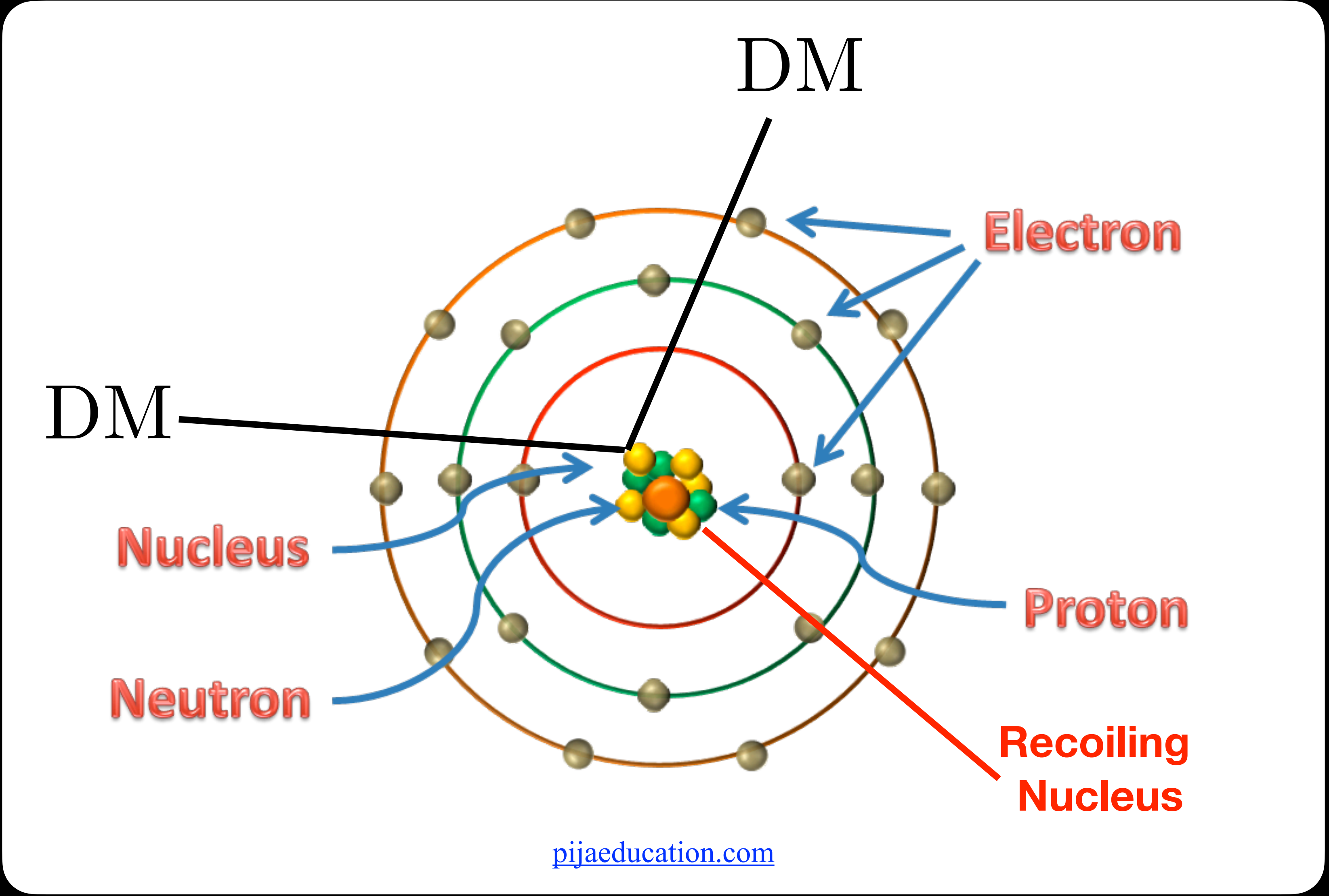
Image: forbes.com



Dark matter hits a nucleus causing a recoil

recoil nucleus is detected and kinematic information used to get dark matter properties

Nuclear Scattering



Nuclear recoil rate

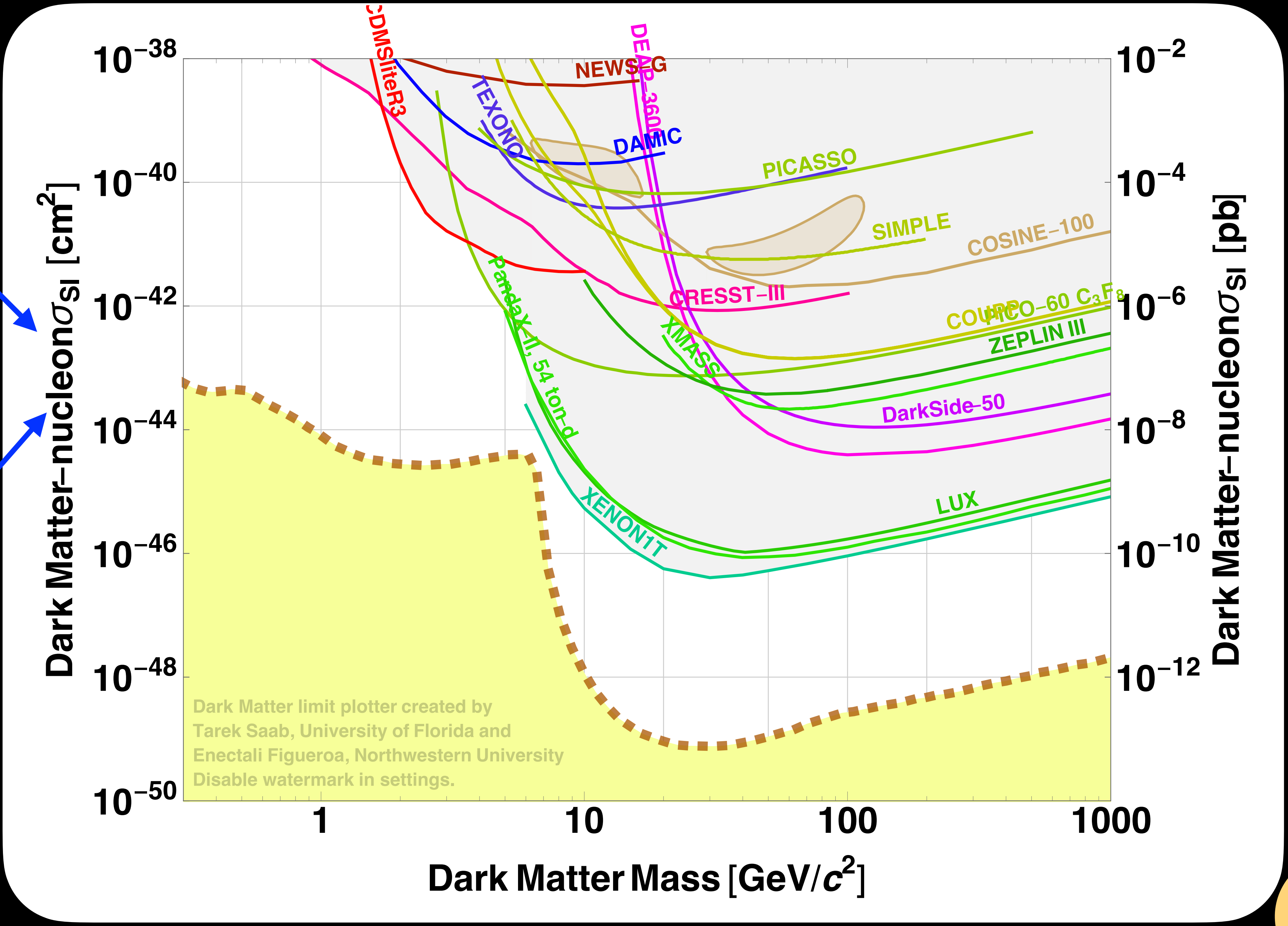
$$\frac{dR}{dE_R} \sim N_T \Delta T \frac{\sigma_{\chi N}}{2\mu^2} |F(E_R)|^2 \frac{\rho_\chi}{m_\chi} \int_{v_{min}}^{\inf} \frac{f(v)}{v} dv$$

Fit recoil rate to experimental data to understand DM parameter space

$$\frac{dR}{dE_R} \sim N_T \Delta T \frac{\sigma_{\chi N}}{2\mu^2} |F(E_R)|^2 \frac{\rho_\chi}{m_\chi} \int_{v_{min}}^{\inf} \frac{f(v)}{v} dv$$

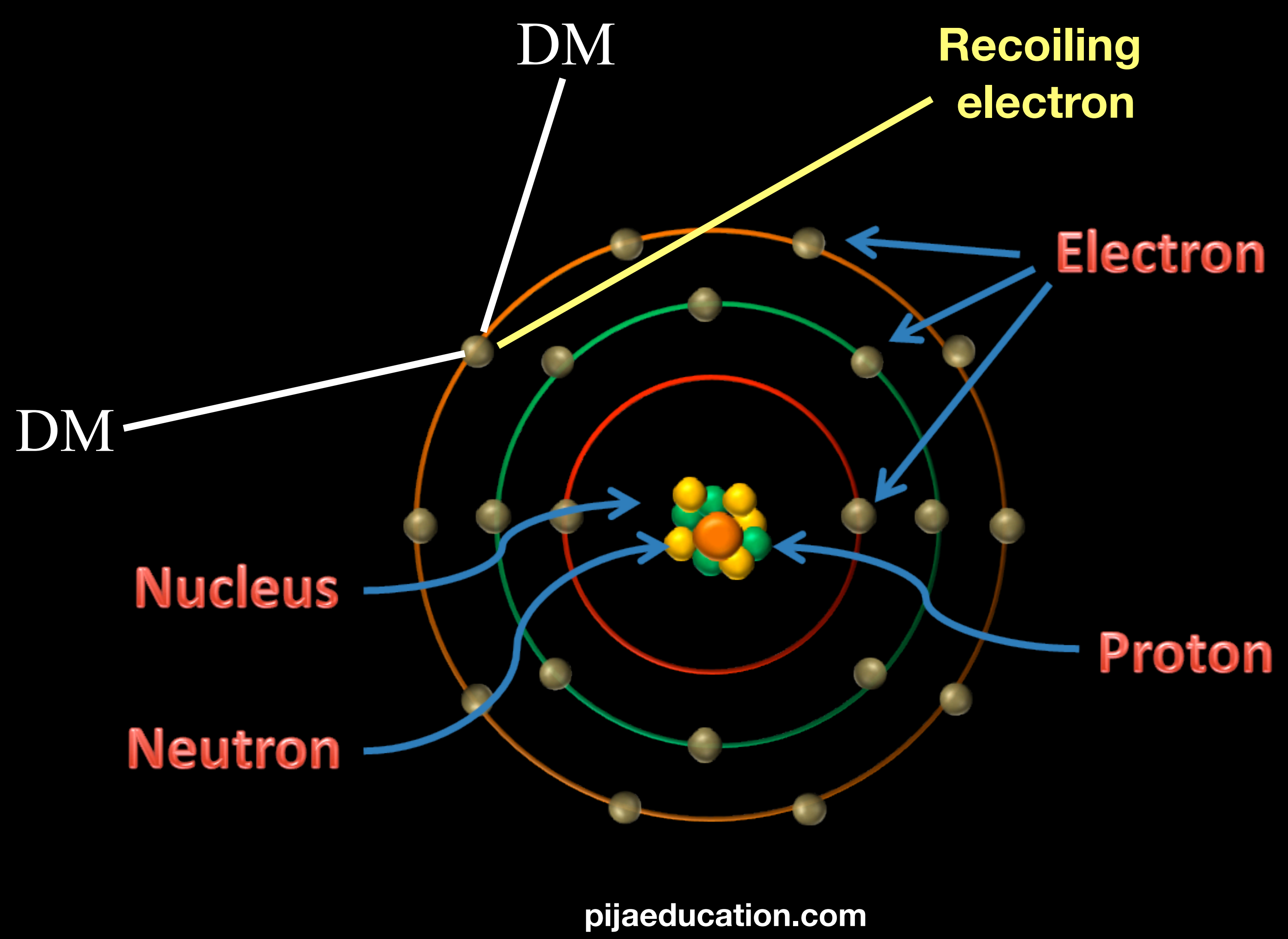
DM-nucleus interaction strength

Particle physics models enter here



supercdms.slac.stanford.edu/dark-matter-limit-plotter

Electron Scattering



Electron recoil rate

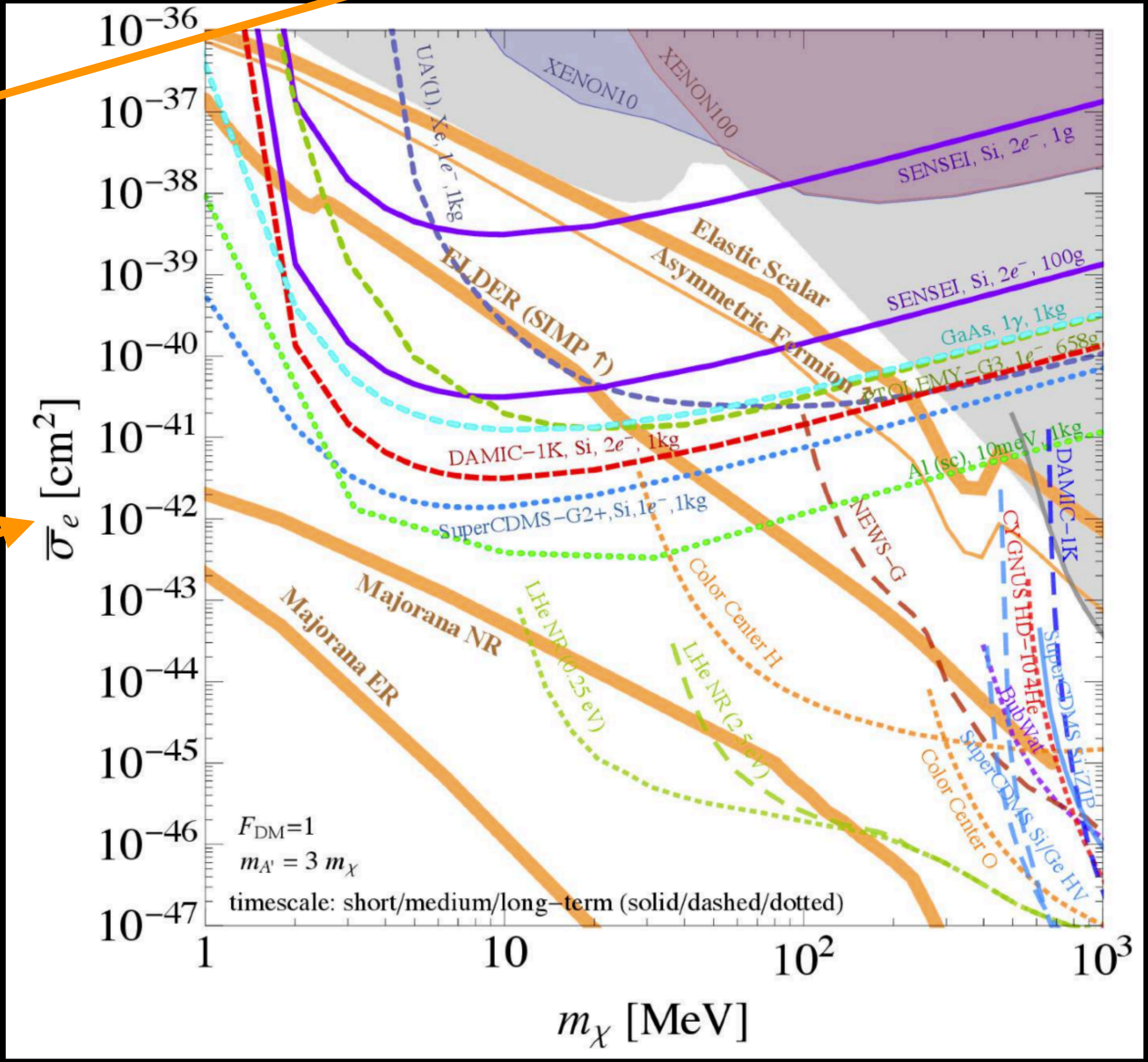
$$\frac{dR}{dlnE_R} \sim N_T \Delta T \frac{\rho_\chi}{m_\chi} \frac{\bar{\sigma}_e}{8\mu^2} \int q dq |f_{ion}^{nl}|^2 |F_{DM}(q)|^2 \int_{v_{min}}^{\inf} \frac{f(v)}{v} dv$$

Fit electron recoil rate to experimental data to understand DM parameter space

$$\frac{dR}{dlnE_R} \sim N_T \Delta T \frac{\rho_\chi}{m_\chi} \frac{\bar{\sigma}_e}{8\mu^2} \int q dq |f_{ion}^{nl}|^2 |F_{DM}(q)|^2 \int_{v_{min}}^{\inf} \frac{f(v)}{v} dv$$

Particle physics information enters here

DM - electron interaction strength



US Cosmic Visions Report: **arXiv:1707.04591**

Example Dark Matter Study

In 2019 I published a paper where I studied Inelastic dark matter at fixed target & collider experiments

PHYSICAL REVIEW D **99**, 115001 (2019)

Revisiting the dark photon explanation of the muon anomalous magnetic moment

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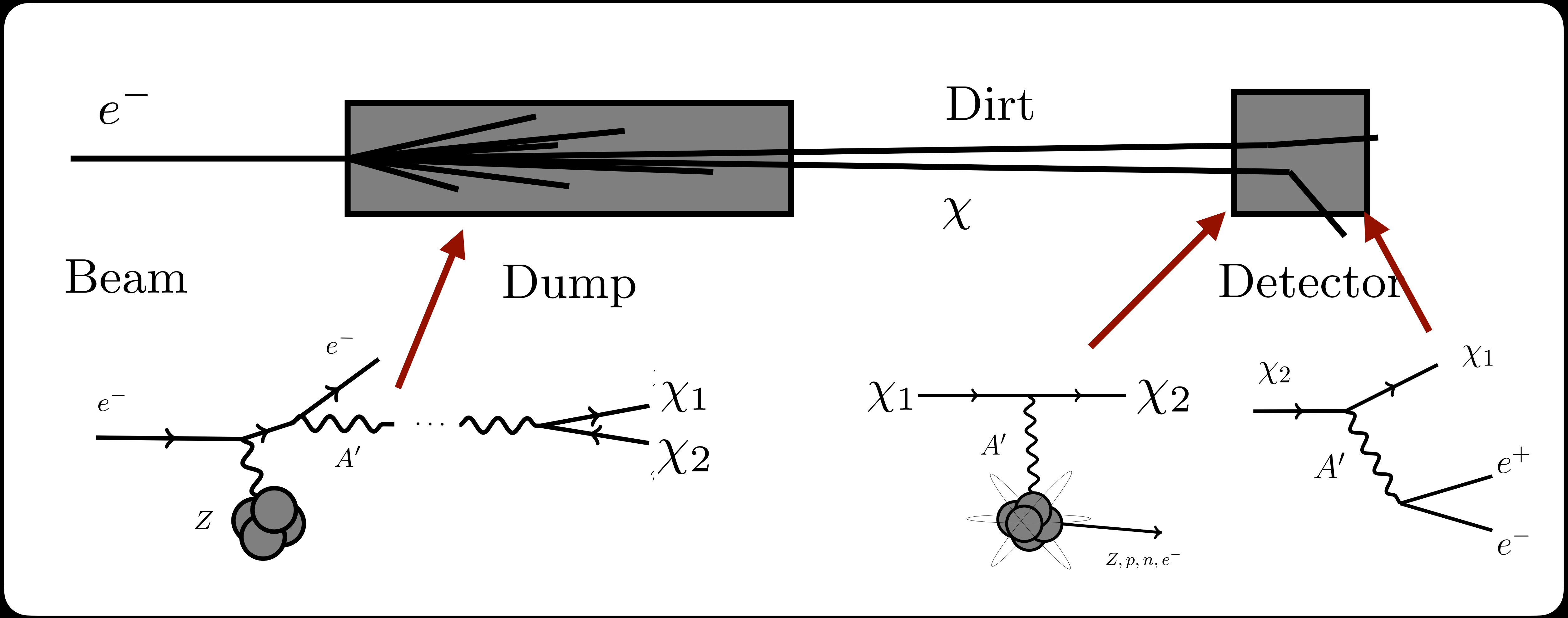
(Received 13 March 2019; published 3 June 2019)

A massive $U(1)'$ gauge boson known as a "dark photon" or A', has long been proposed as a potential explanation for the discrepancy observed between the experimental measurement and theoretical determination of the anomalous magnetic moment of the muon ($g_\mu - 2$) anomaly. Recently, experimental results have excluded this possibility for a dark photon exhibiting exclusively visible or invisible decays. In this work, we revisit this idea and consider a model where A' couples inelastically to dark matter and an excited dark sector state, leading to a more exotic decay topology we refer to as a semivisible decay. We show that for large mass splittings between the dark sector states this decay mode is enhanced, weakening the previous invisibly decaying dark photon bounds. As a consequence, A' resolves the $g_\mu - 2$ anomaly in a region of parameter

DM is accompanied by heavier dark sector particle

Dark sector connected to SM via new force carrier called dark photon

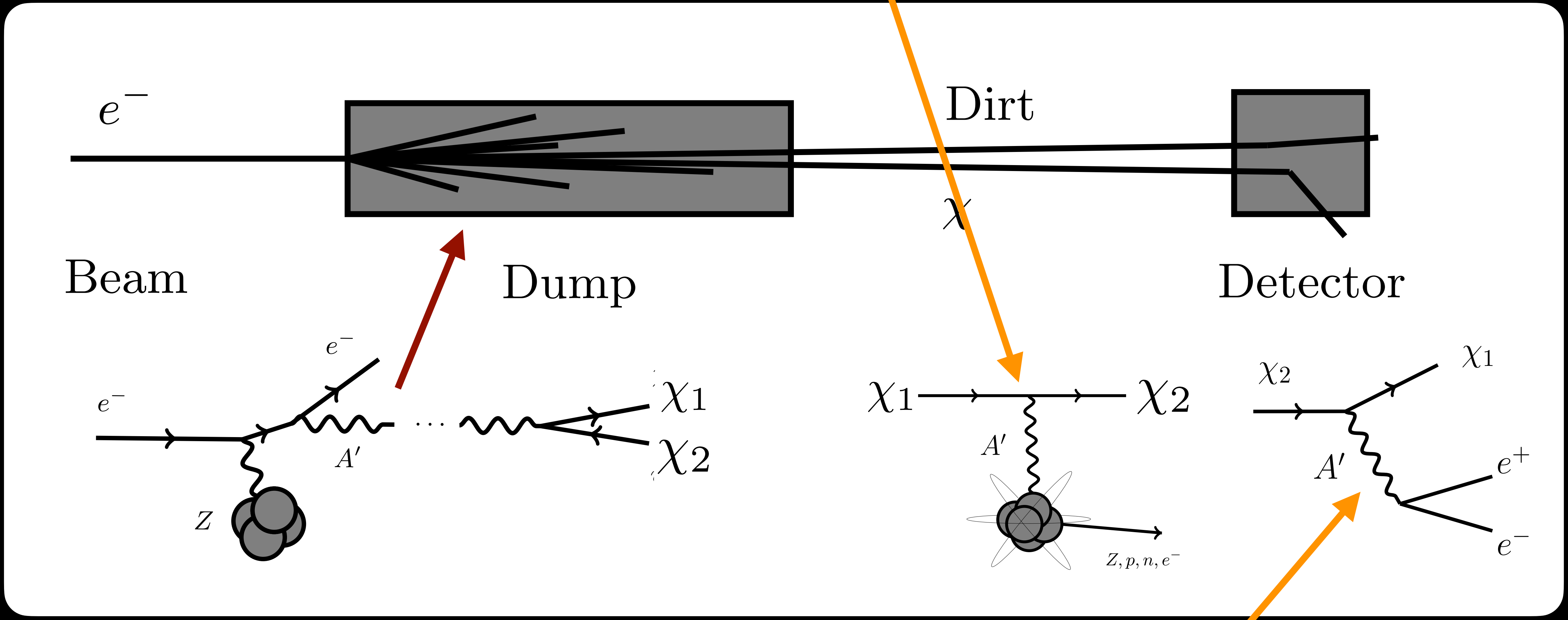
Dark photon is produced in proton & electron fixed target experiments and decays to inelastic dark matter inside experiment



Dark photon is produced in fixed target experiments and decays to inelastic dark matter inside experiment

DM and heavier dark sector state travel to detector

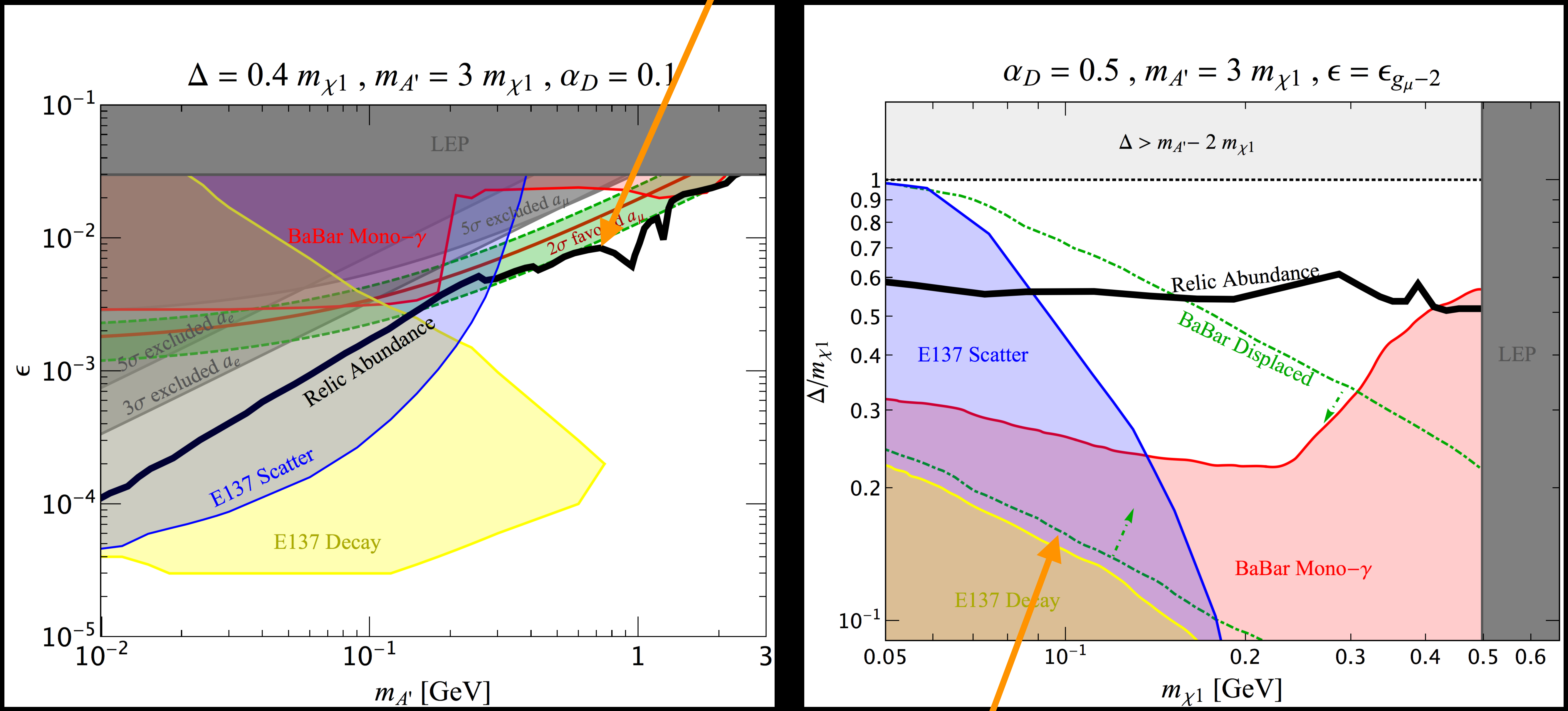
- DM can scatter with SM inside detector



- heavy state can decay to DM and SM particles inside detector

Experimental constraints on this model

Relic density calculation compared to data



Theory computations compared to accelerator experiment data

Recap: lecture 2

- Cosmological abundance and production of dark matter

- Calculation and Understanding of Relic density from Freeze-out

- Methods of detecting dark matter

Indirect detection

Production at colliders/Accelerators

Dark matter direct detection

- Example Inelastic dark matter study

Questions?

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