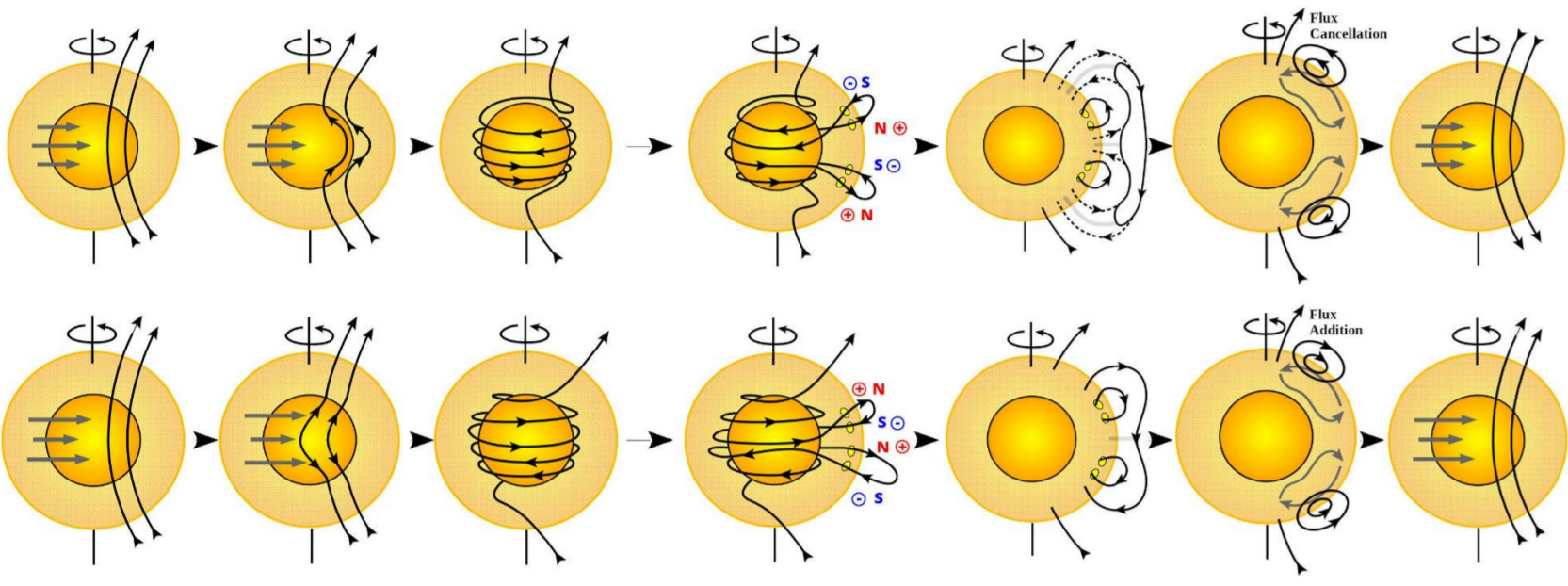
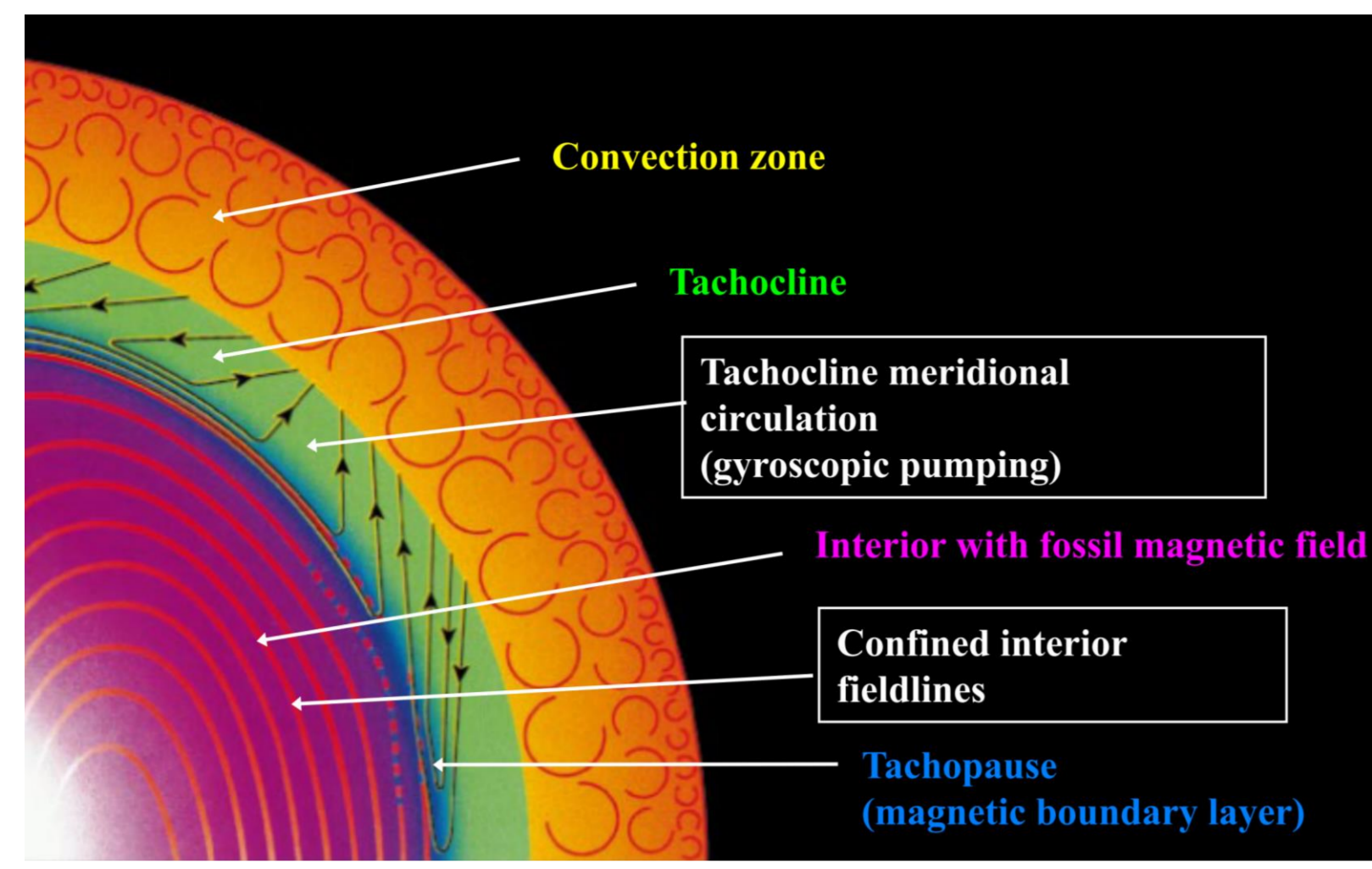


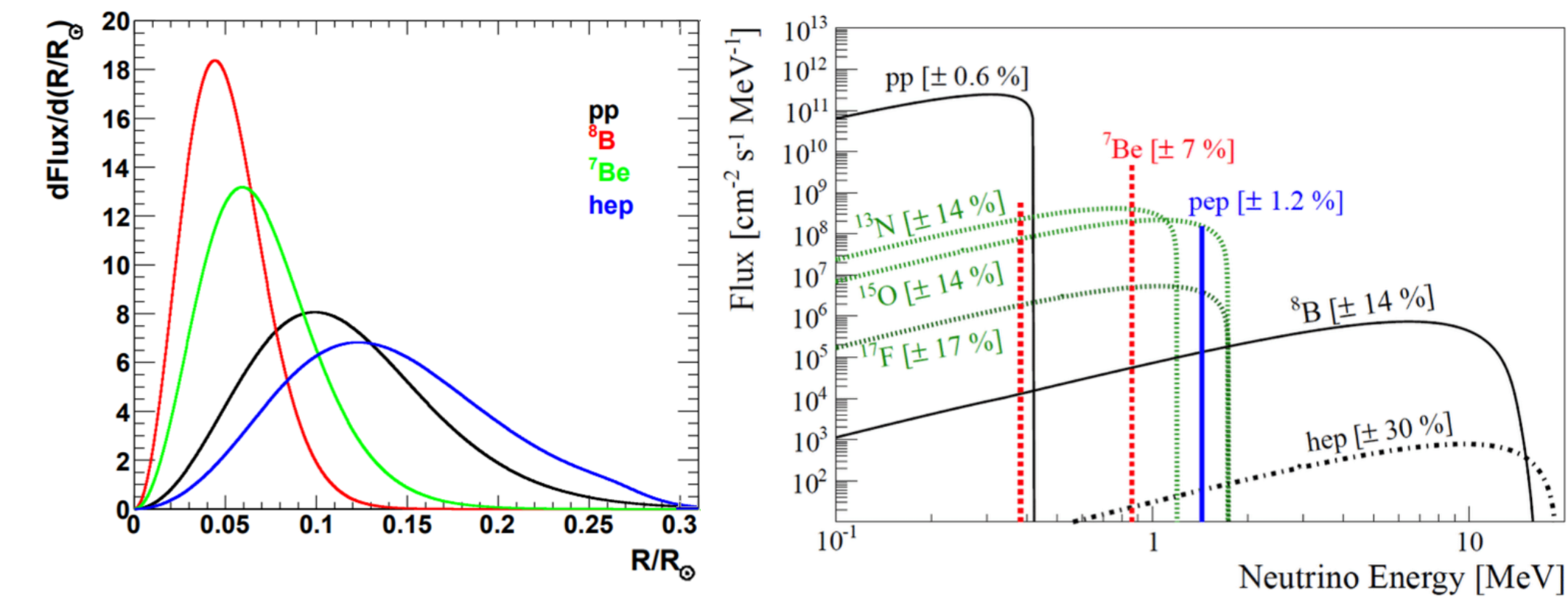
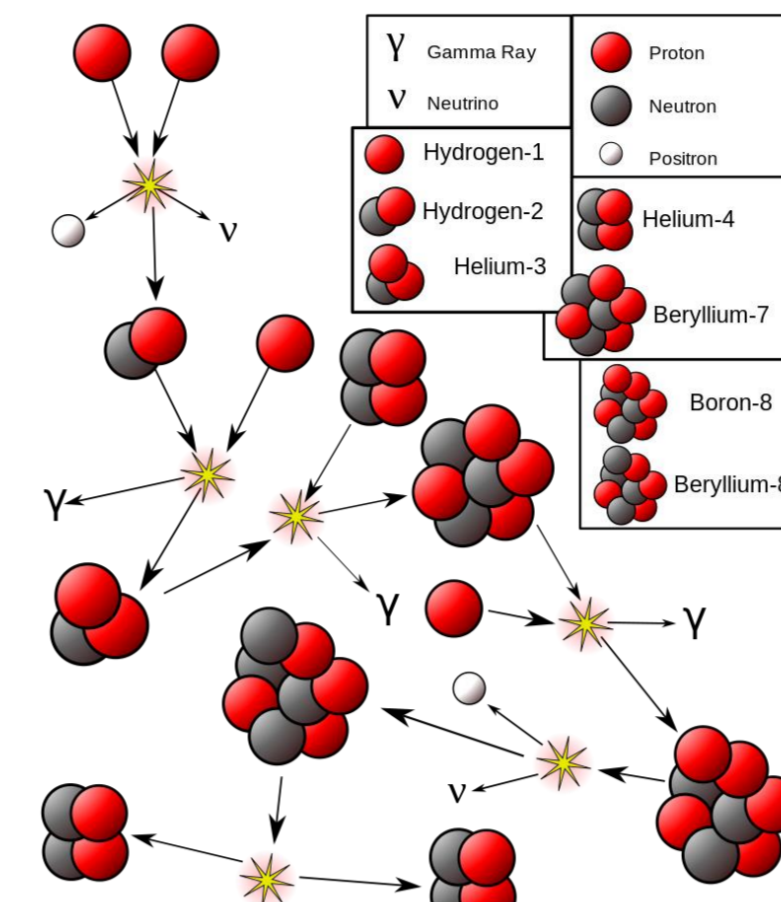
Motivation



- LMA solution
 - $1/r^2$ annual modulation
- Flux dependence of core temperature

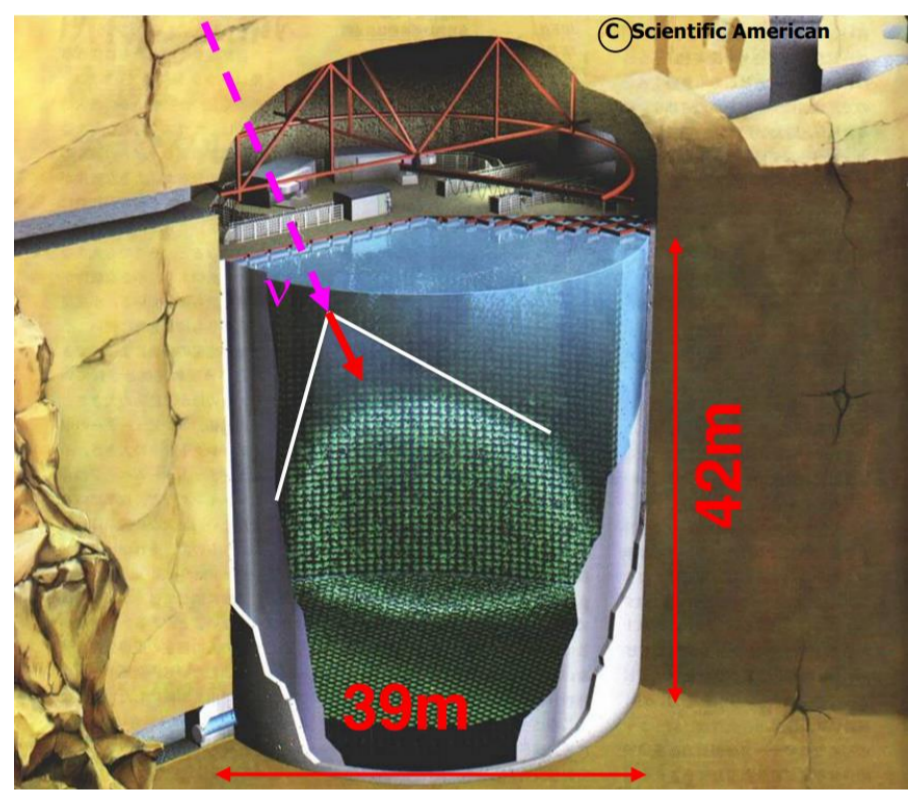
$$\Phi_{\text{solar } 8B \nu} \propto T^{25}$$
 (PRD, 53:4202, 1996)
- Friction in tachocline
 - Related with the sunspot
 - Strong magnetic field
- Magnetic field flips flavor and spin of ν called resonant spin flavor precession
 - Additional disappearance
 - Periodicity of flux modulation

Solar neutrino at SK

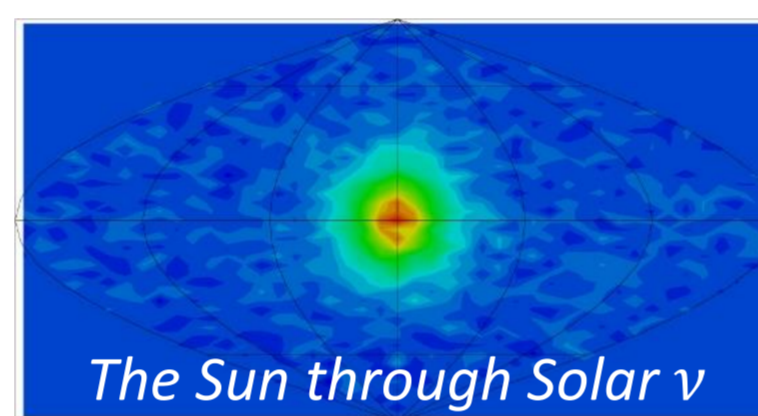


- Neutrino from 8B(99.8%) and hep(0.2%)
- Most neutrino made in core ($<0.3R_{\text{sun}}$)
- During escaping the Sun, ν_{solar}
 - MSW effect inside of the Sun
 - Coincide 2nd mass eigenstate
- Measure elastic scattering electron/muon neutrino event at SK
 - $\nu_e + e \rightarrow \nu_e + e$ (Z, W) / $\nu_\mu + e \rightarrow \nu_\mu + e$ (Z)

Super-Kamiokande (SK)



- Water cherenkov detector
- Water 50ktons (22.5ktons fid.)
- 1,000m underground
- Inner-Detector : 11,146 50cm PMTs (40%)
- Outer-Detector : 1,885 20cm PMTs



SK Phase	Start date ~ End date	Live days	Energy range [MeV]	Flux (ϕ_ν)+(stat.)+(sys.) [$10^6 \text{cm}^{-2} \text{s}^{-1}$]
SK-I	1996-05-31 ~ 2001-07-15	1495.7	4.49~19.5	$2.35 \pm 0.02 \pm 0.08$
SK-II	2002-12-10 ~ 2005-10-06	791.9	6.49~19.5	$2.38 \pm 0.05^{+0.16}_{-0.15}$
SK-III	2006-05-23 ~ 2008-08-17	548.5	4.49~19.5	$2.32 \pm 0.04 \pm 0.05$
SK-IV	2008-09-15 ~ 2018-05-30	2967.7	4.49~19.5	$2.31 \pm 0.014 \pm 0.040$

Annual modulation

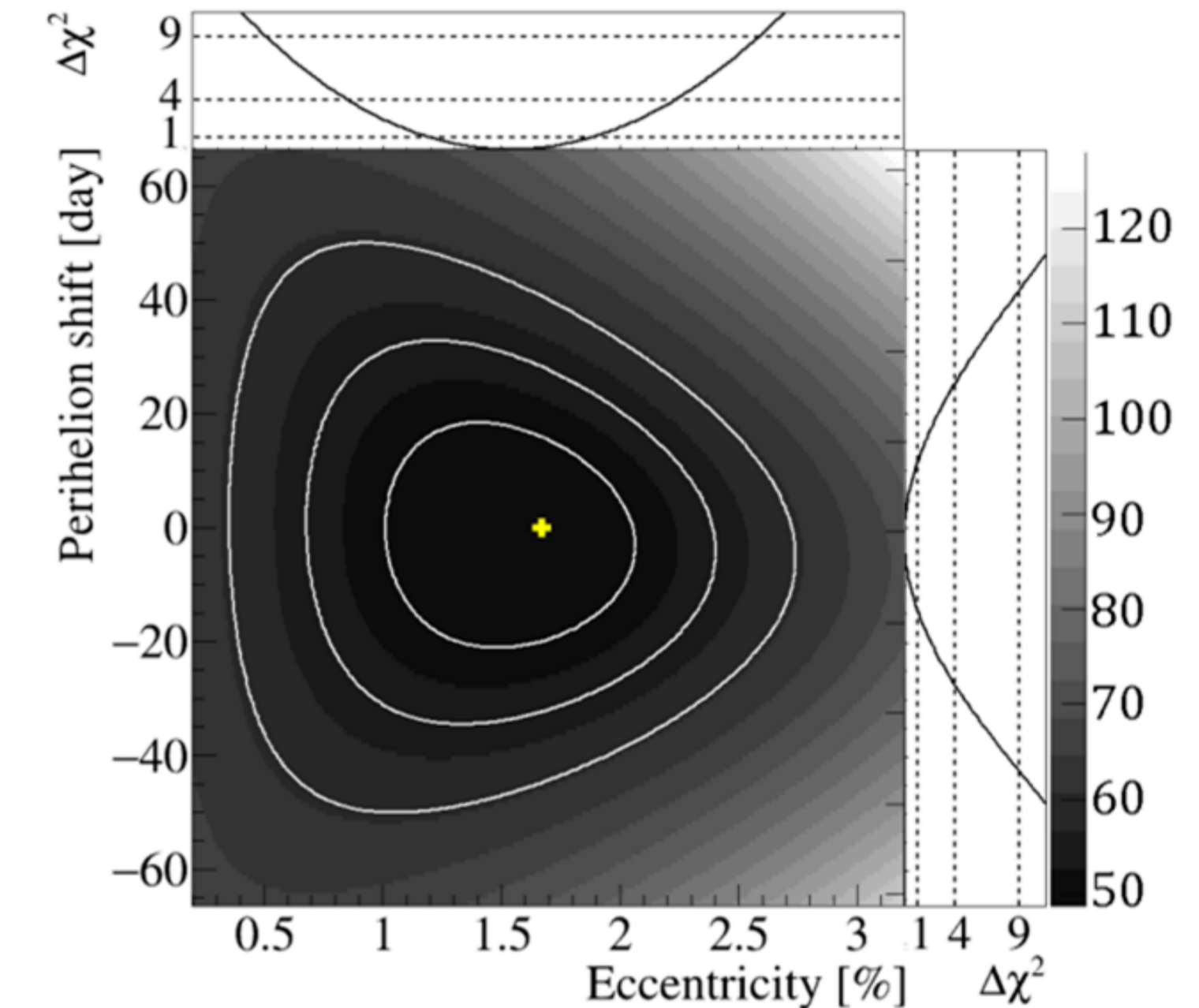
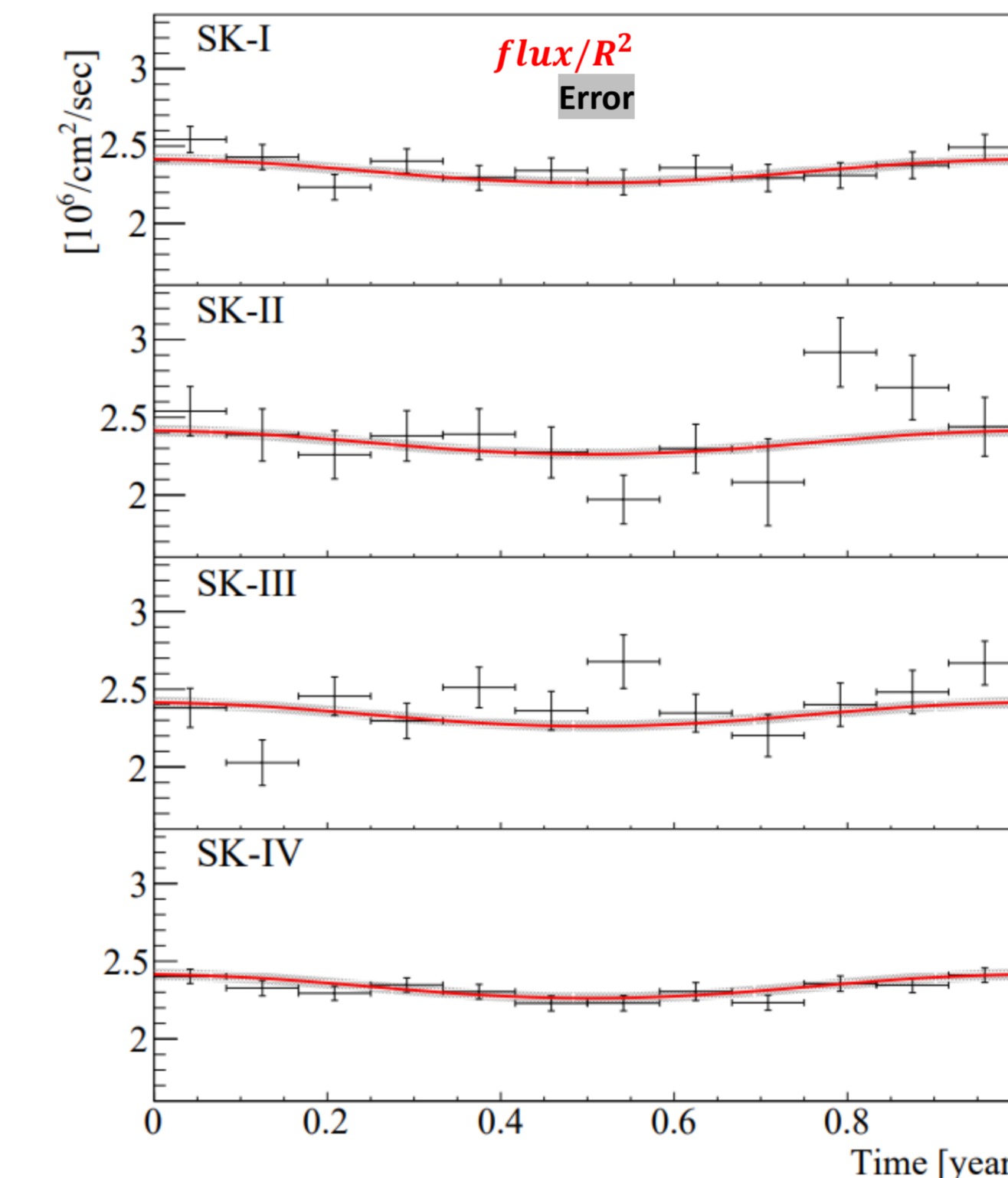
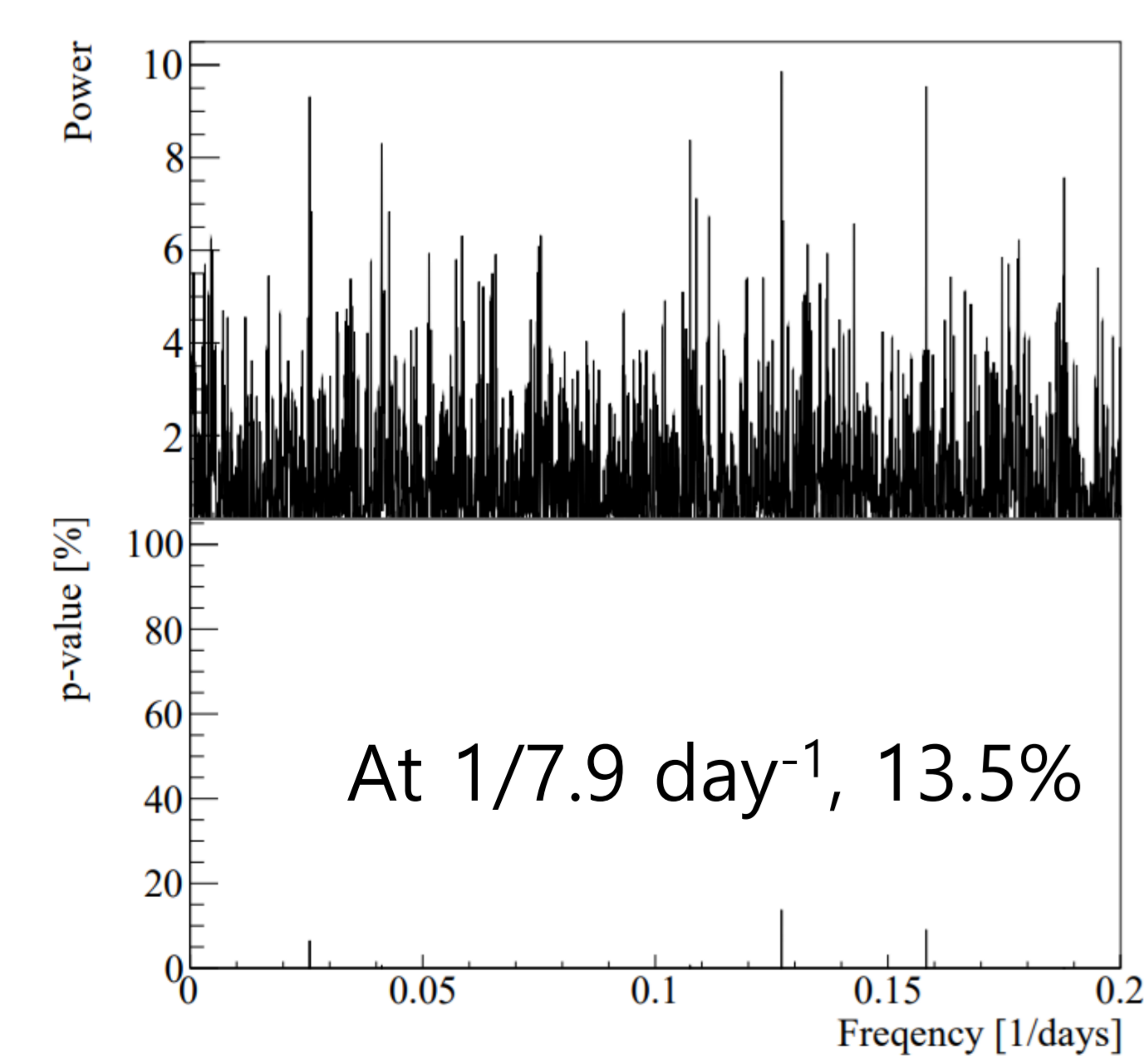
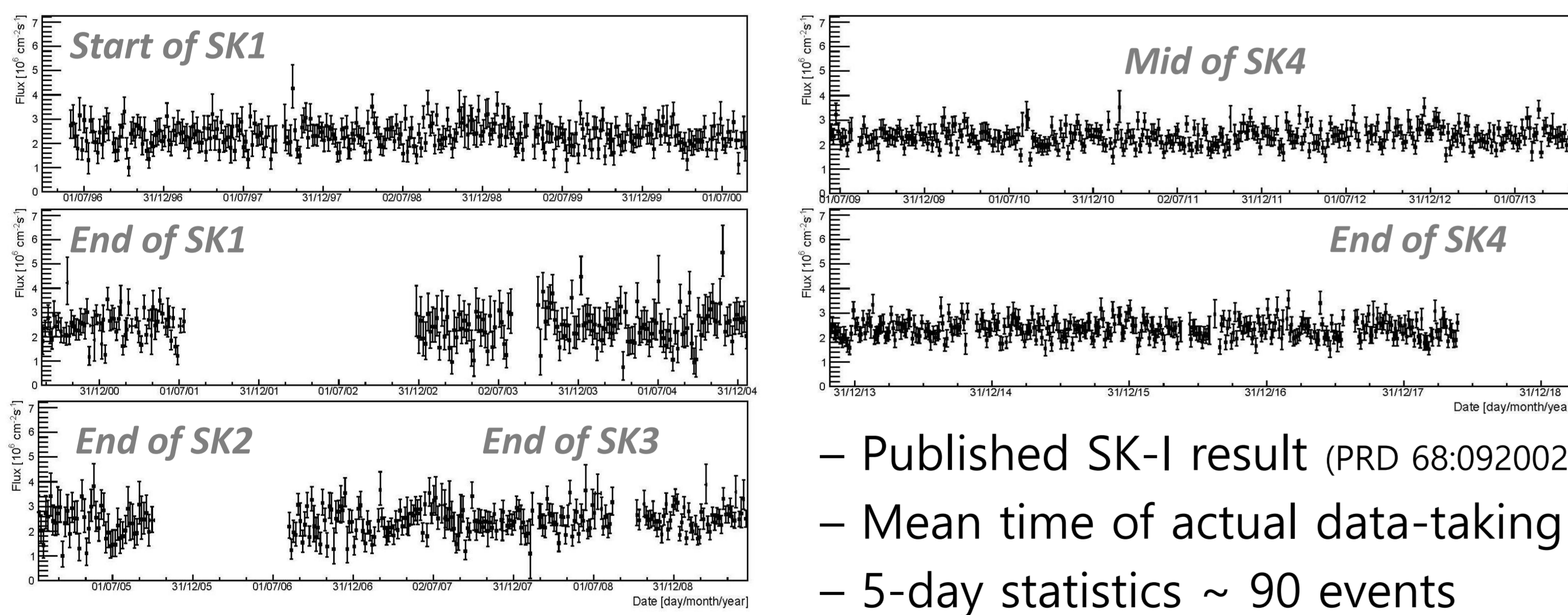


FIG. 5. $\Delta\chi^2$ for Kepler constants (ϵ , t_p) with contours $\Delta\chi^2=2.3, 6.18, 11.83$; $\epsilon = 1.53 \pm 0.35\%$ (known as 1.67%), $\delta t_{\text{peri}} = -1.5 \pm 13.5$ days shifted from Jan 3rd with minimum $\chi^2/NDF = 48.74/(48-3)$.

- The seasonal variation of the solar neutrino flux for SK-I/II/III/IV.
- Best fit flux at 1 AU $f = (2.335 \pm 0.036) \times 10^6 \text{cm}^{-2} \text{s}^{-1}$
- Kepler constants in measurement are consistent to the expected.

Search for Periodicity



- Average flux g_0 with statistical asymmetric error (σ_r)
- Minimize $-\ln L$.

$$-\ln L(r; A, B) = \frac{1}{2} \sum_r \left(\frac{D_r - g(r)}{\sigma_r} \right)^2$$

$$g(t_i, t_f) = g_0(t_i, t_f) + \frac{1}{\Delta t} \int_{t_i}^{t_f} A \cos(\omega t) + B \sin(\omega t) dt$$
- Power = $\ln \left(\frac{L(g(A, B))}{L(g_0)} \right)$
- For p-value, generate 10,000 MC samples for g_0 with asym. error.

\rightarrow No significant periodicity for >5 days

Sensitivity

- $1/R^2$ corrected 5-day data
- Simulated 1,000 MC signals
- Amplitude / null flux : 1-99%
- Frequency : 1000 cases
- Use asymmetric stat. error

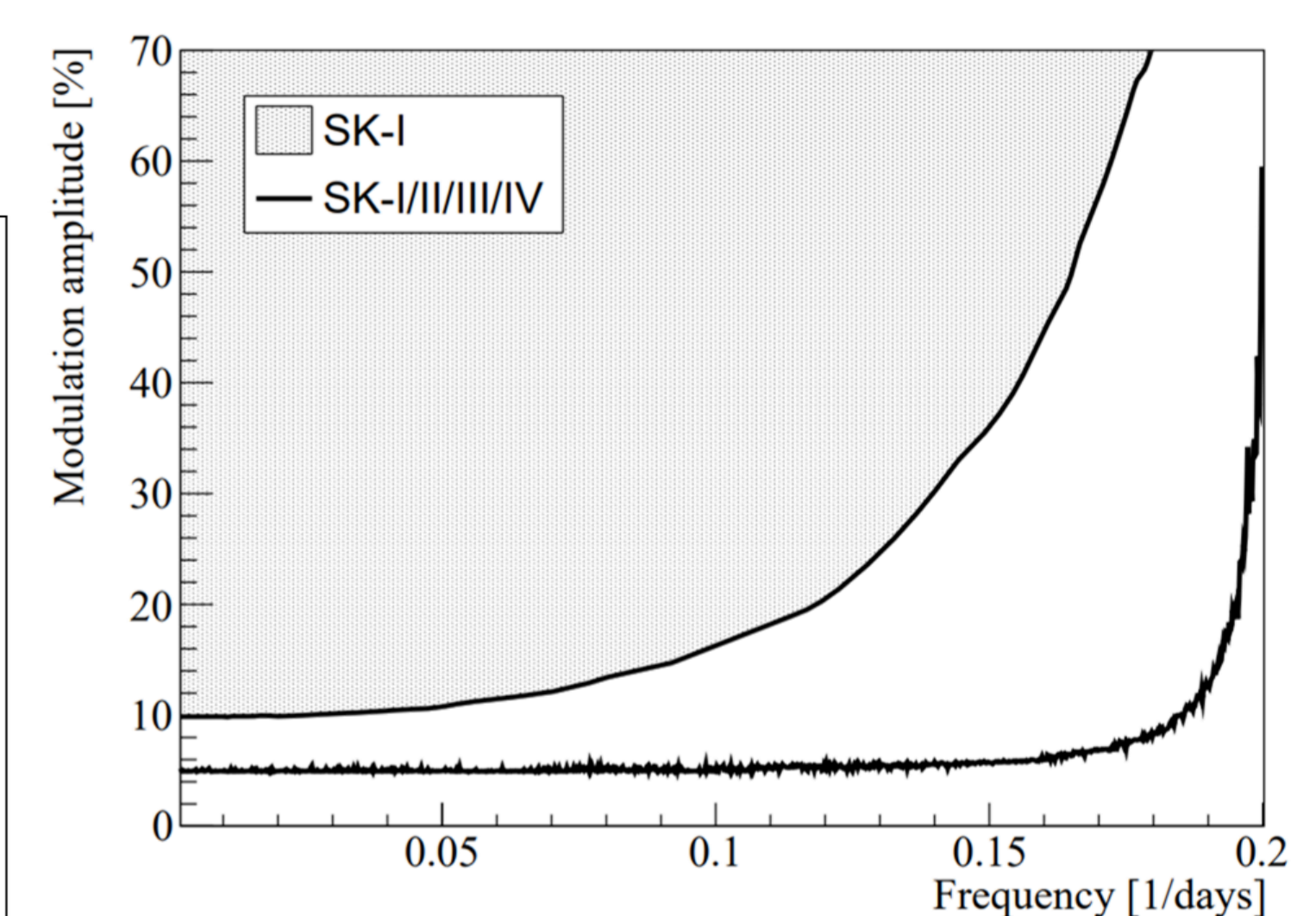
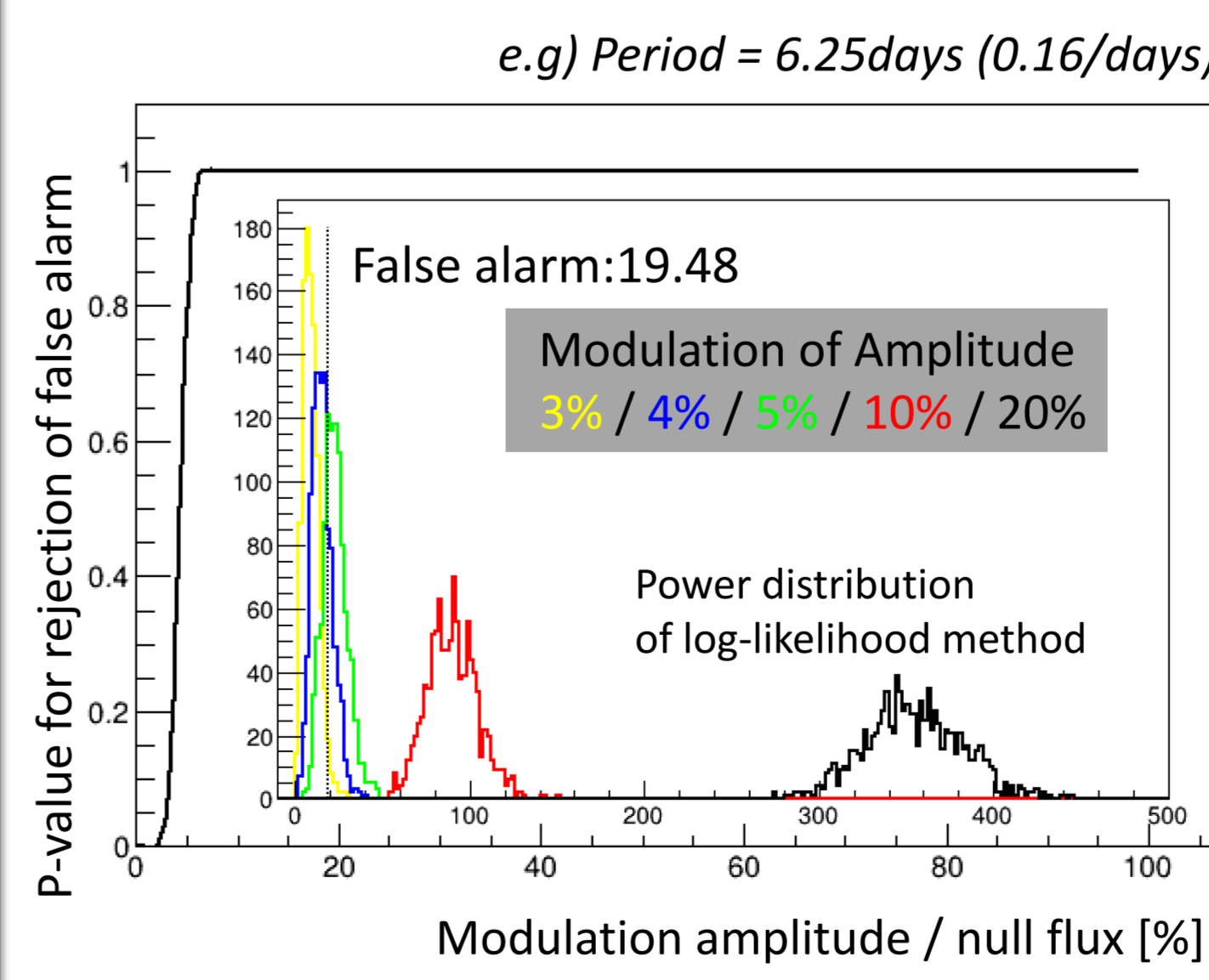


FIG. 3. A search sensitivity of solar neutrino modulation amplitude for each frequency with CL 95%. The shaded area is ruled out using SK-I solar neutrino data. The solid curve shows this analysis results.

Rule out the solar modulation amplitude greater than 5.1%
 \rightarrow Significant improvement compared to SK-I results

Summary

- Measure Solar neutrino flux using SK-I/II/III/IV data.
- Confirm the annual modulation with measurement of Kepler constants as expected.
- For the periodicity longer than 5-day interval, we do not find significant peak.
- We can rule out the modulation of 5.1% amplitude or larger amplitude (period > 6.25 day).