Searching for Dark Matter with the RES-NOVA detector

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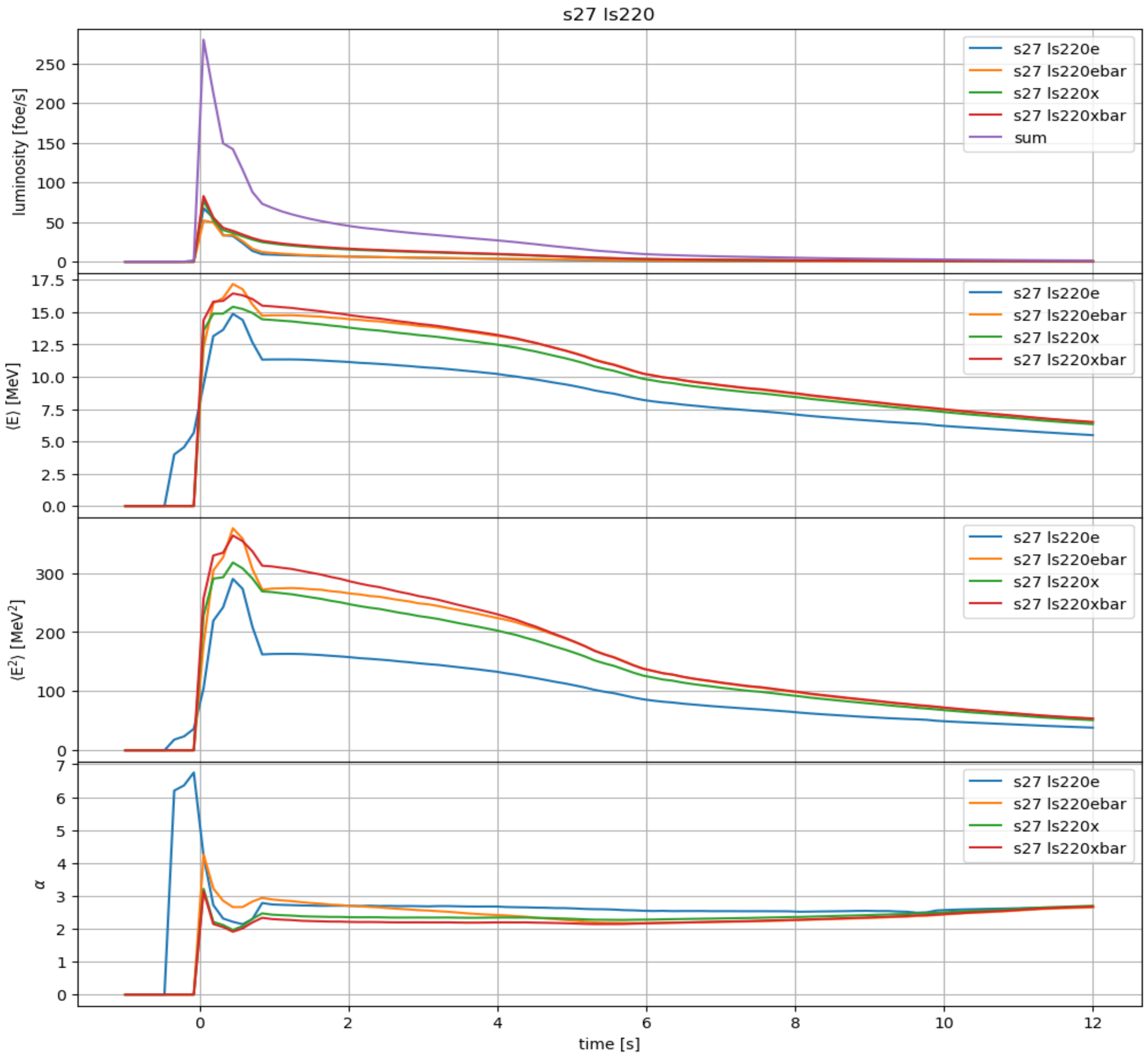
RES-NOVA

From latin -> “New thing”

Latin -> more on that later

NOVA reminds to Supernovae



The neutrino source

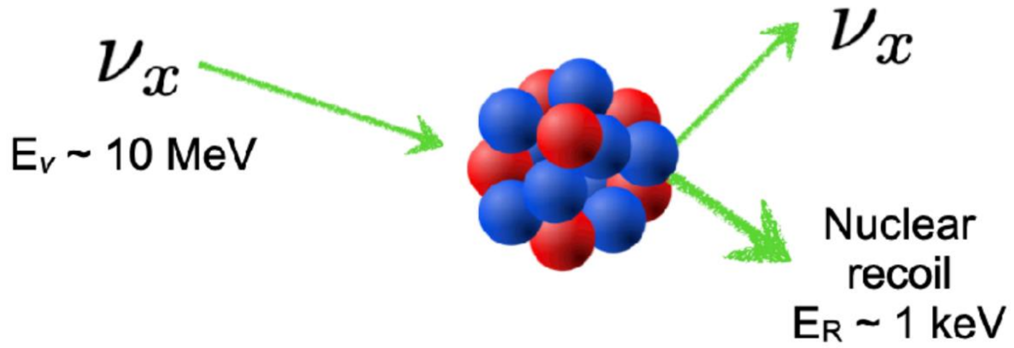
$$f^0_\beta(E,t) = \frac{L_\beta(t)}{4\pi d^2}\frac{\phi_\beta(E,t)}{\langle E_\beta(t)\rangle}$$

$$\phi_\beta(E,t) = \xi_\beta(t)\left(\frac{E}{\langle E_\beta(t)\rangle}\right)^{\alpha_\beta(t)}$$

$$\exp\left(-\frac{(\alpha_\beta(t)+1)E}{\langle E_\beta(t)\rangle}\right)$$

Coherent Elastic ν Nucleus Scattering

$$\frac{d\sigma}{dE_R} = \frac{G_F^2 m_N}{8\pi(\hbar c)^4}\left[(4\sin^2\theta_W - 1)Z + N\right]^2 \left(2 - \frac{E_R m_N}{E^2}\right) \cdot |F(q)|^2 \, ,$$



> Equally sensitive to all ν-flavors

> High interaction cross-section

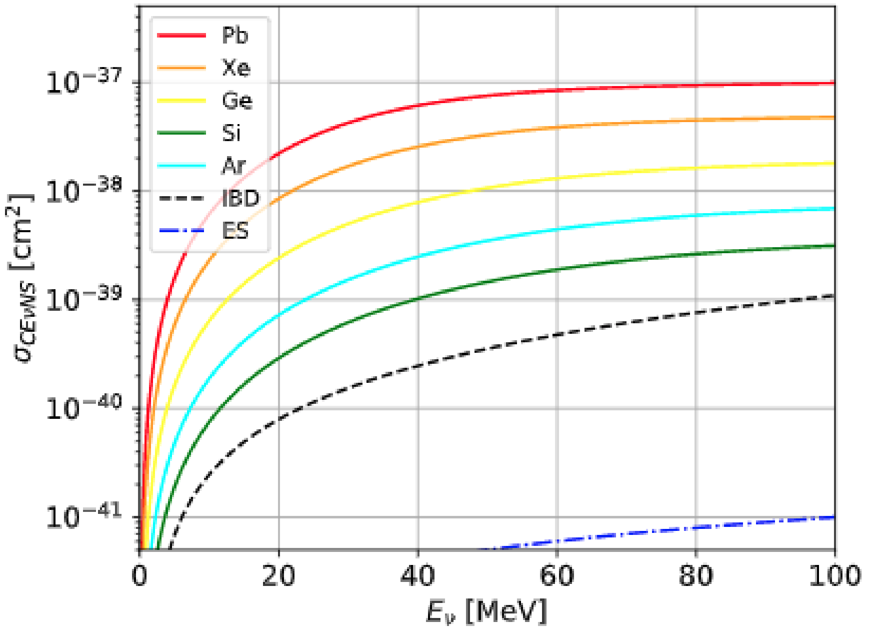
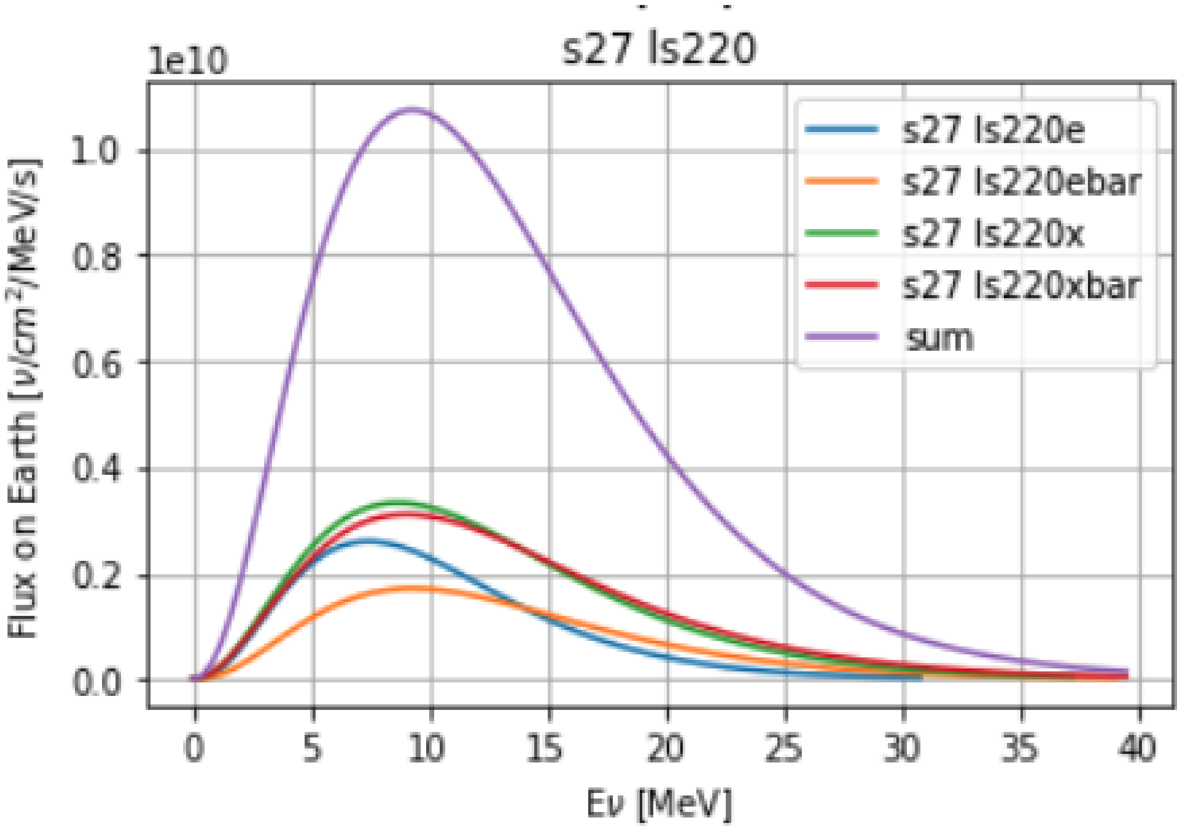


FIG. 2. Coherent elastic neutrino-nucleus scattering (CEνNS) cross sections as a function of the energy of the incoming neutrino for different target nuclei. The dashed lines show the inverse-beta decay (IBD) and neutrino elastic scattering on electrons (ES) cross-sections for comparison. Given the high cross-section, CEνNS has the potential to provide large statistics with small detector volumes.

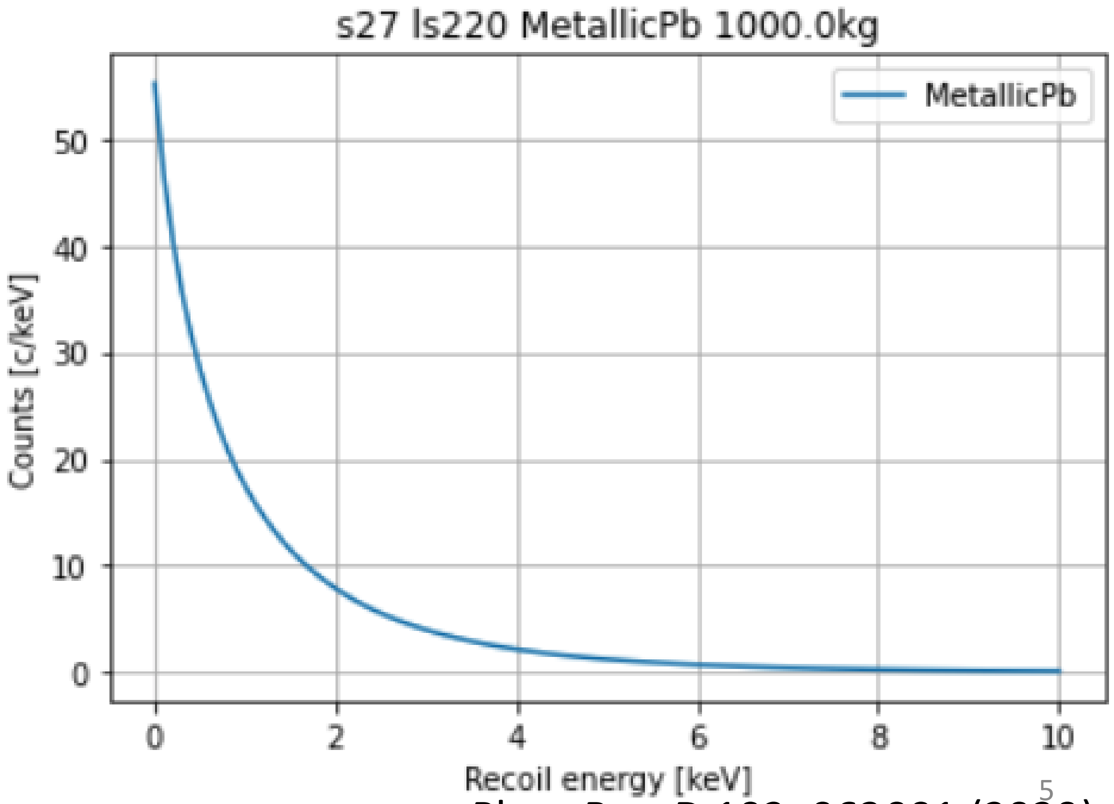
Phys. Rev. D 102, 063001 (2020)

SN CEνNS in Pb Target (on Earth)

The emitted neutrino spectrum is (almost) Maxwell-Boltzmann



Observed nuclear recoil spectrum



Phys. Rev. D 102, 063001 (2020)

From SN neutrinos to Dark Matter

- Pb can proble N^2 weak-scattering down to $10^{-38}/10^{-39}cm^2$
- (Spin-independent WIMPs) Dark matter couples to A^2 and obeys a Maxwell-Boltzmann distribution

$$\frac{d\Gamma}{dE_R} = \frac{\rho_\chi}{m_N m_\chi} \int_{v_{min}}^{\infty} d^3\overline{v} f(\overline{v}) v \frac{d\sigma(v, E_R)}{dE_R}$$

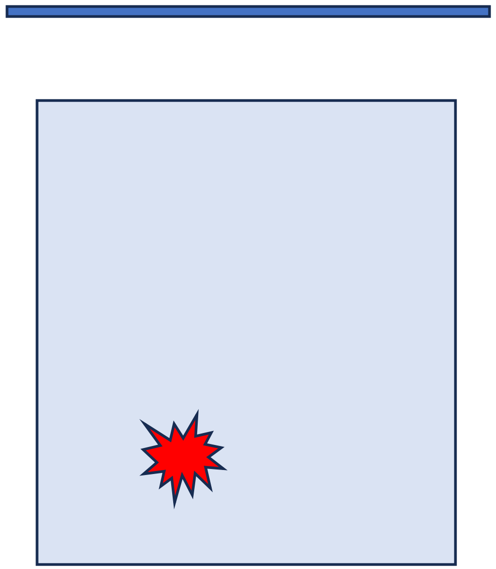
$$\left(\frac{d\sigma}{dE_R}\right)_{SI} = \frac{m_N \sigma_0}{2\mu^2 v^2} F^2(E_R)$$

$$\sigma_0 = \frac{4}{\pi} A^2 f^2 \mu^2$$

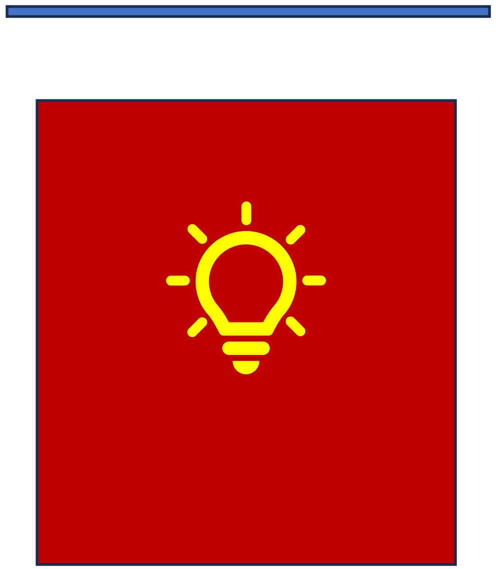
- Dark matter is a always-on source, making its flux 10^7 greater than SN neutrinos (for a 1y measurement).

A detailed calculation yields that we can reach 10^{-43} cm^2 of sensitivity

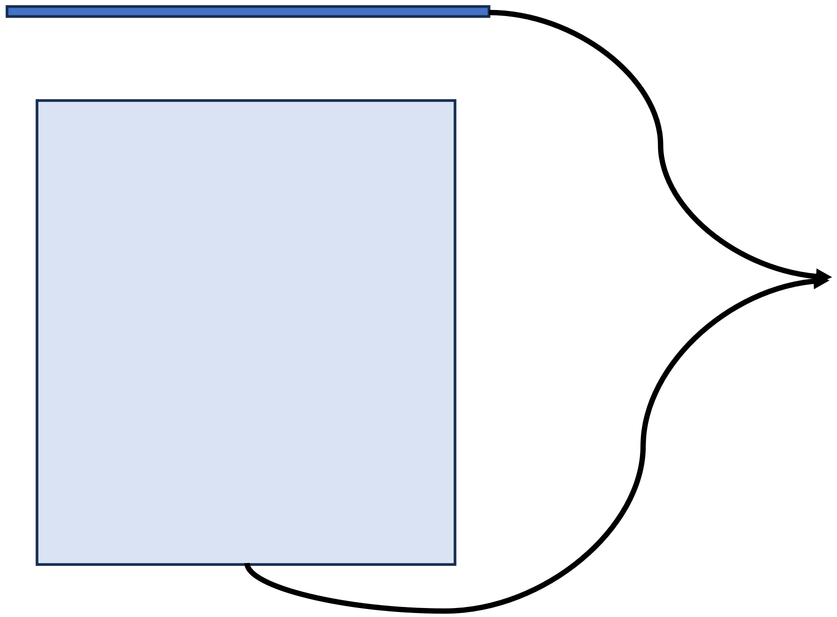
The RES-NOVA detector

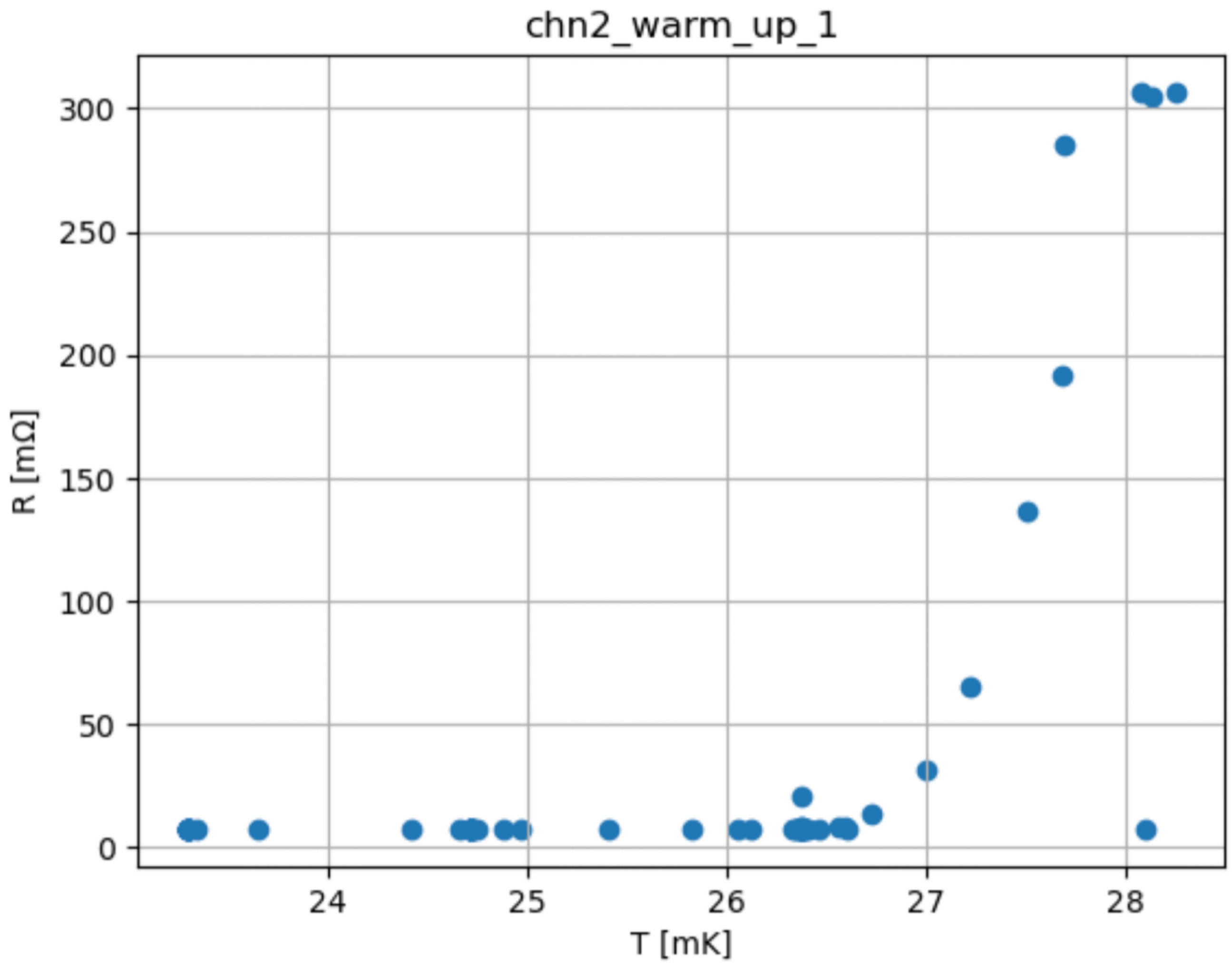


The RES-NOVA detector

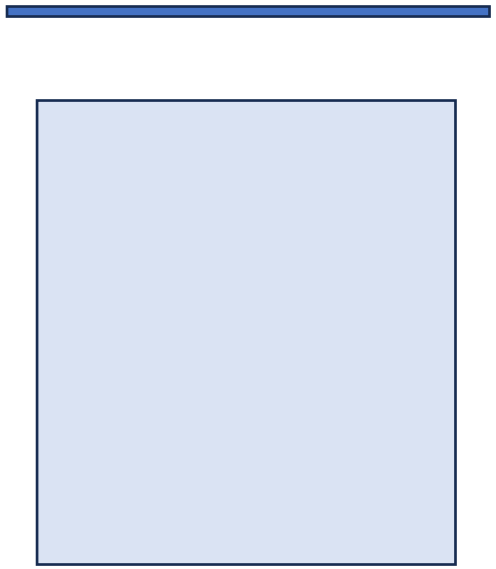


The RES-NOVA detector

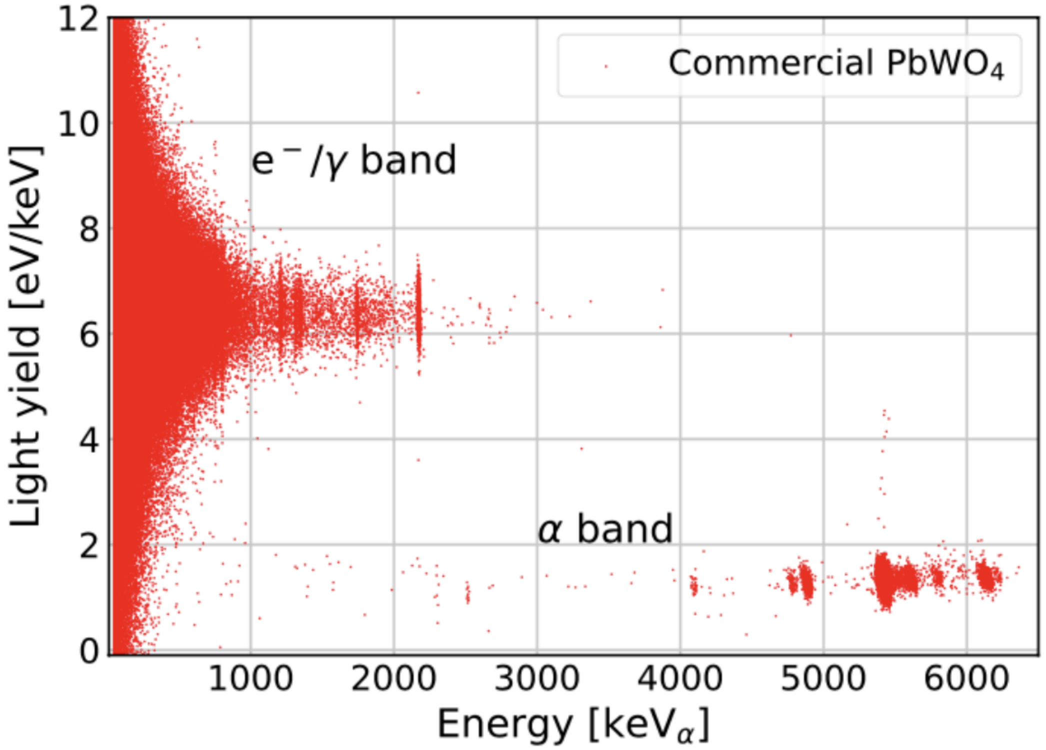




The RES-NOVA detector



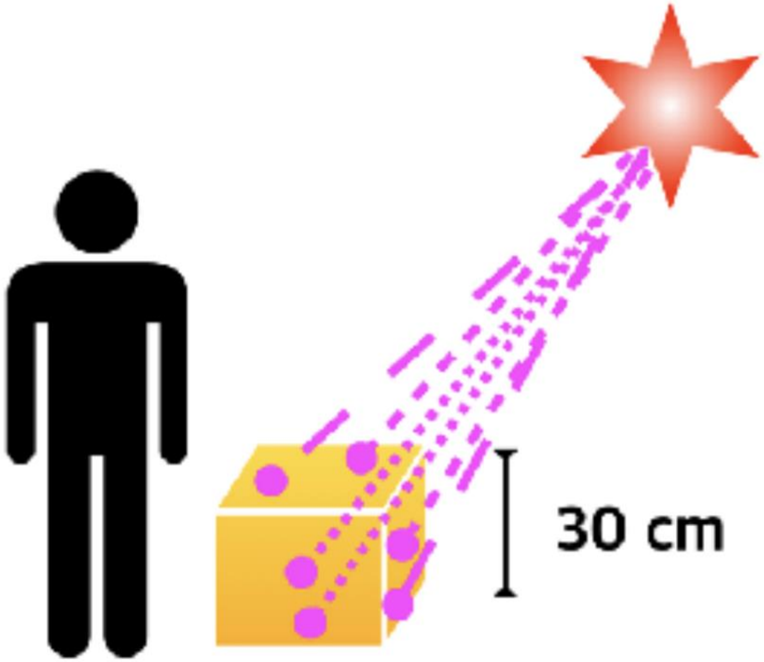
Detector energy spectrum of a cryo-$PbWO_4$

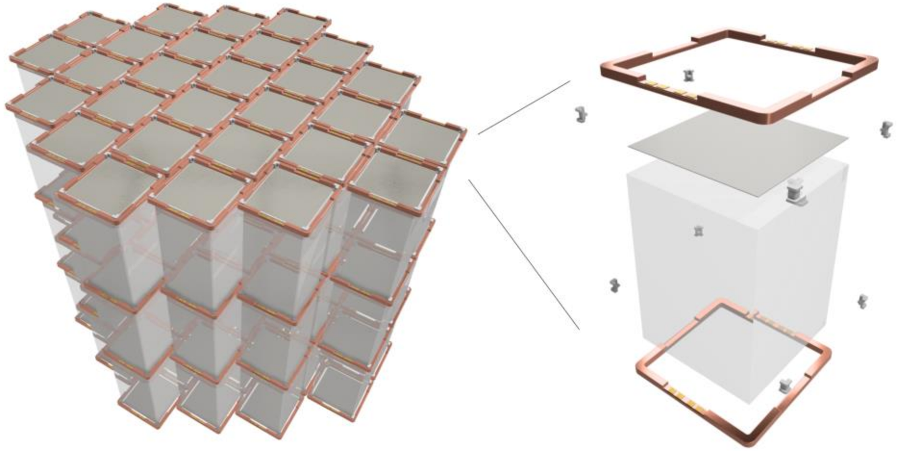


J.W. Beeman, **LP** et al., Eur. Phys. J. A 49, 50 (2013)

The RES-NOVA detector

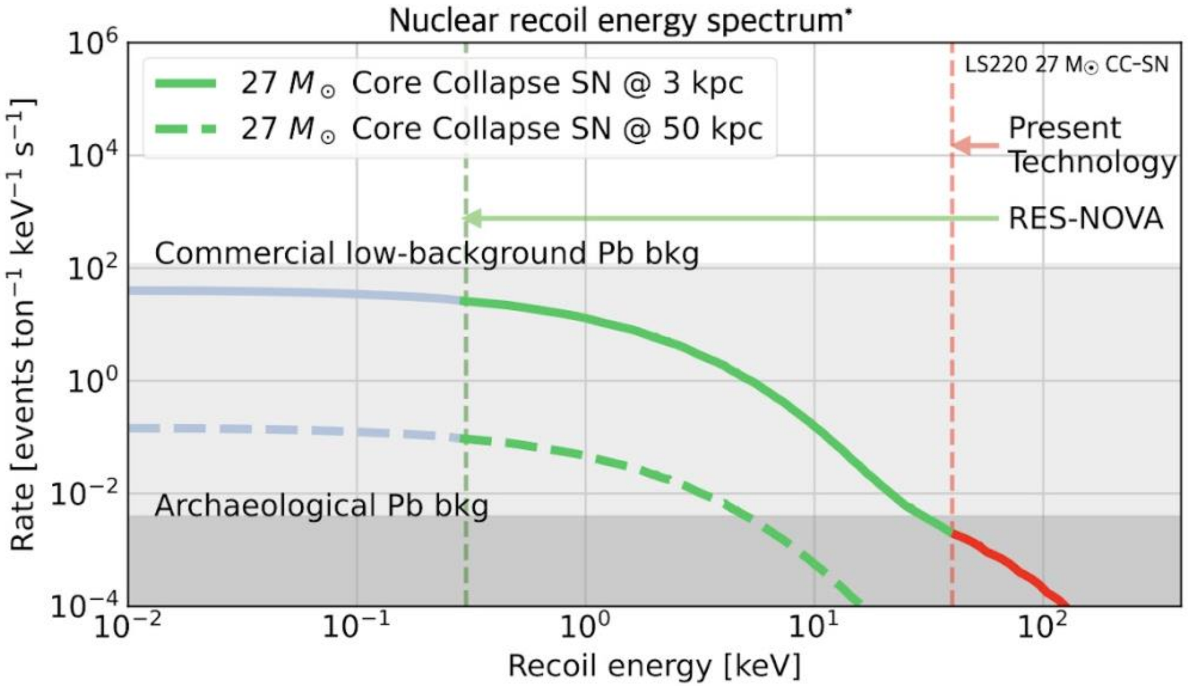
- Array of $PbWO_4$ crystals operated as (scintillating) cryogenic detectors ($8.28\ g\cdot cm^{-3}$)
- Scintillating cryogenic detectors provide powerful background rejection thanks to the simulaneous read-out of phonon and light channels. Time coincident analysis of different detector modules allows for further background suppression
- Energy measured by means of sensitive Transition Edge Sensors (1sigma resolution: 200 eV)
- TESs have already demonstrated the capability of sub-keV nuclear recoil energy threshold





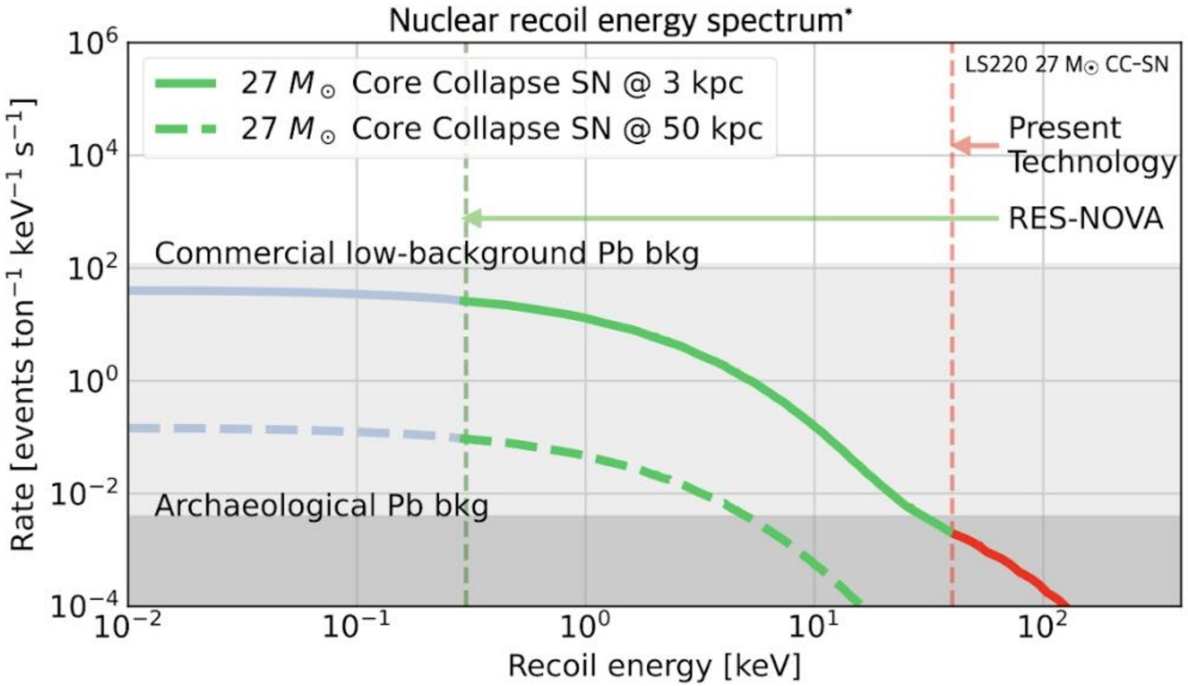
The downfall of Pb

Commercial Pb has 10^4 Bq/ton of radioactive ^{210}Pb (Q-value 63 keV, $\tau_{1/2}$=22 y). That's bummer!



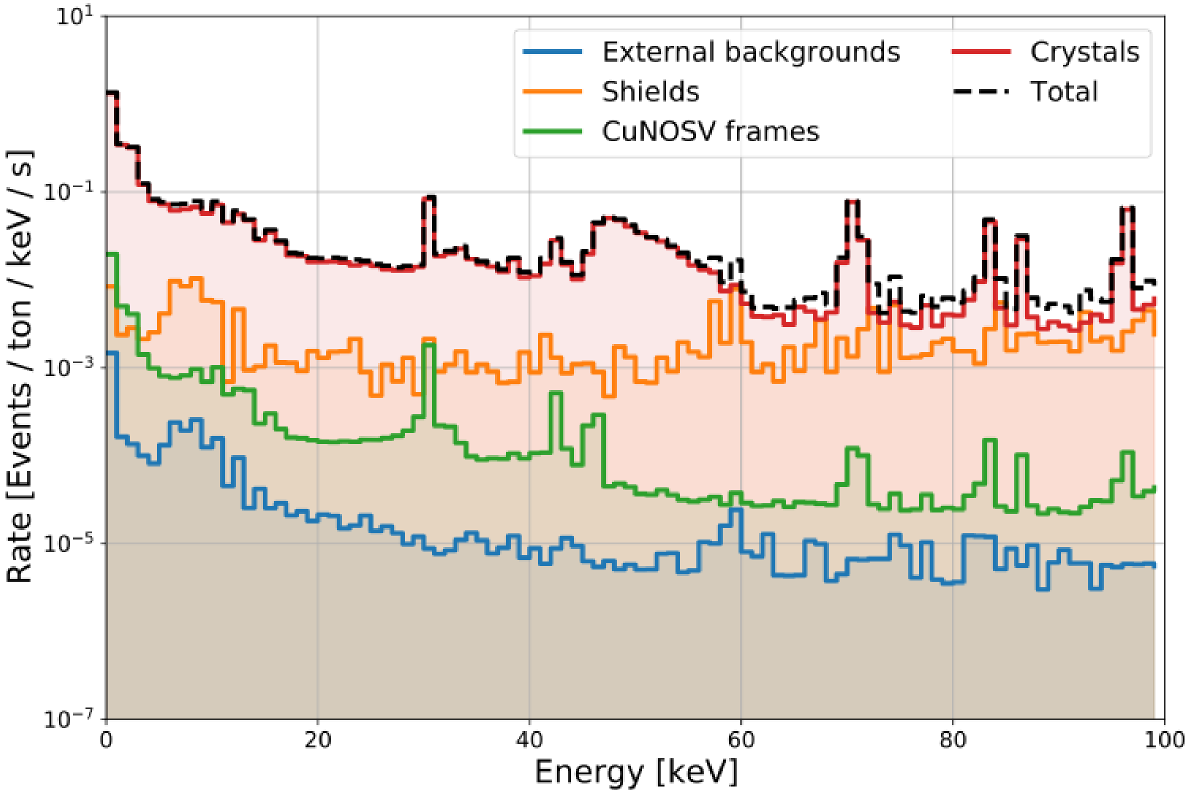
And the solution

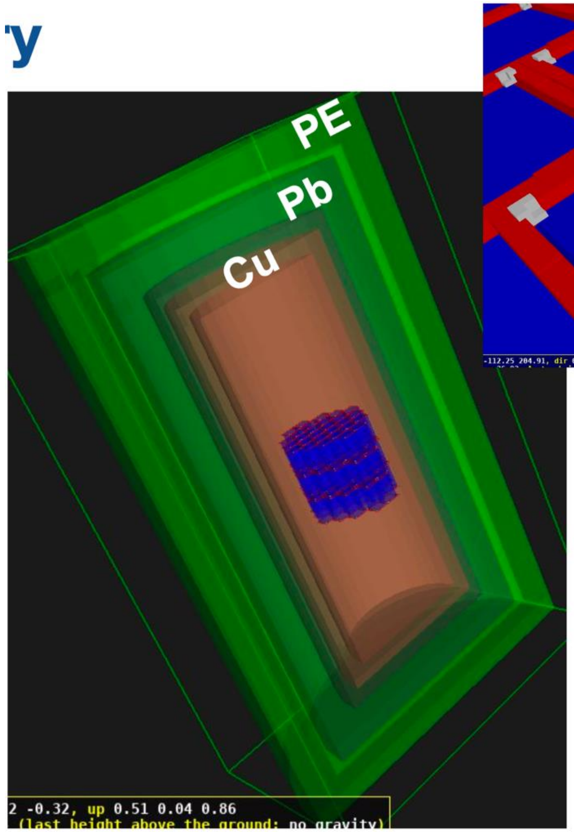
Commercial Pb has 10^4 Bq/ton of radioactive ^{210}Pb (Q-value 63 keV, $\tau_{1/2}$=22 y). That’s bummer!

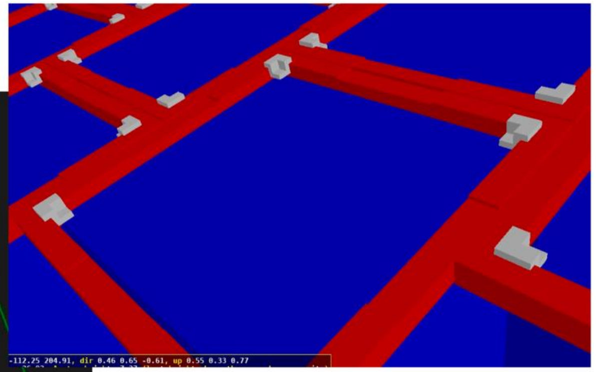


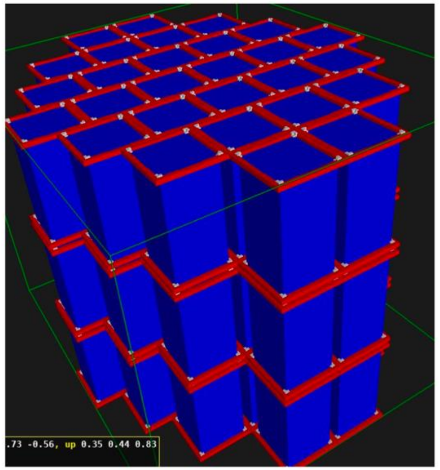
We will deploy $PbWO_4$ grown with 2000years old archaeological lead ^{210}Pb is expected to be below 1mBq/kg

Our background model - complete

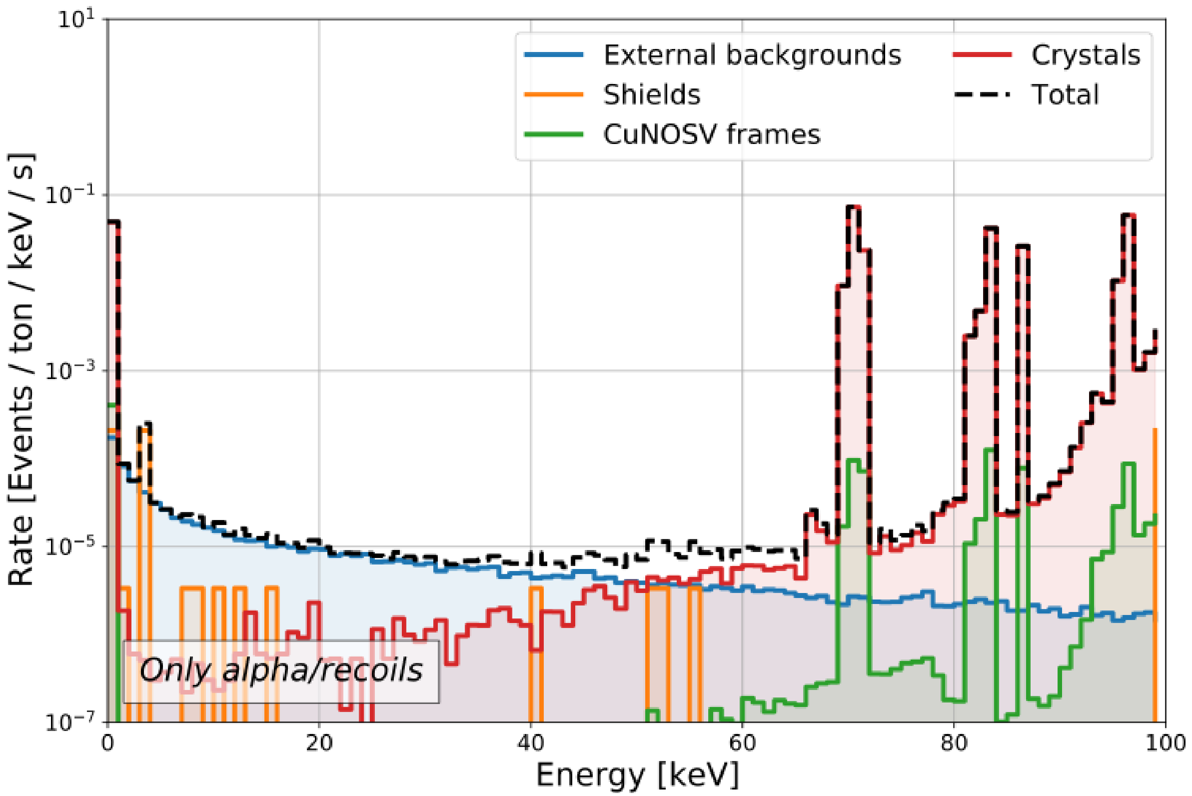


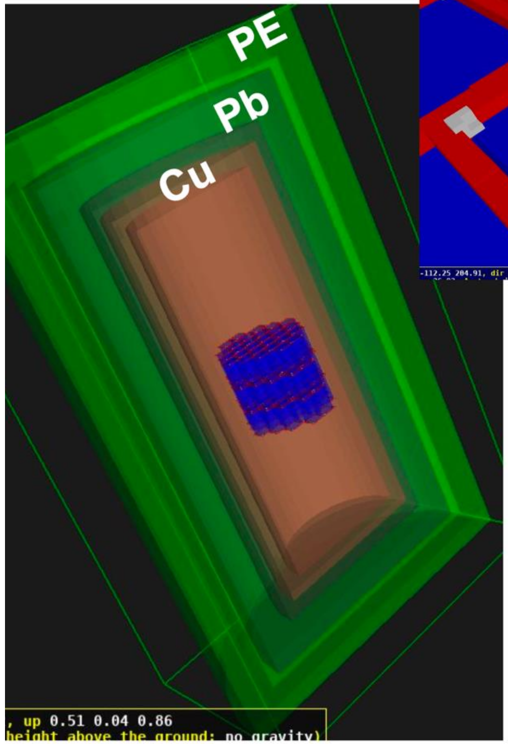


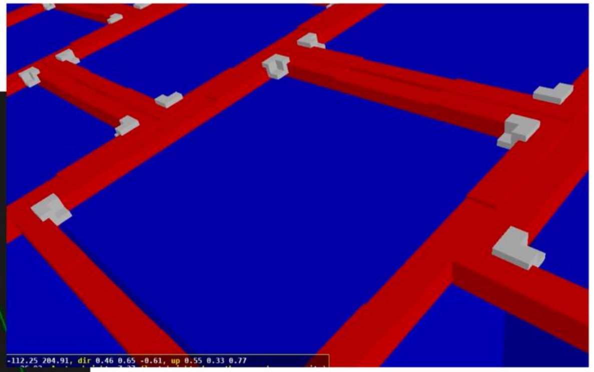


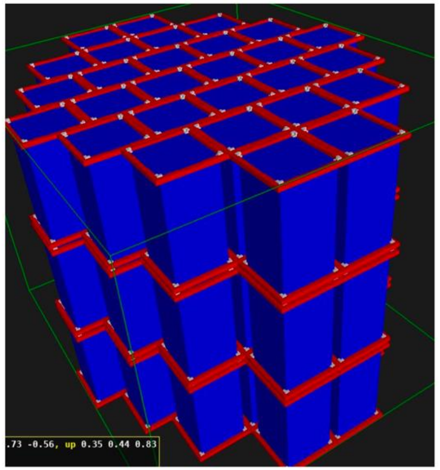


Our background model – nuclear recoils





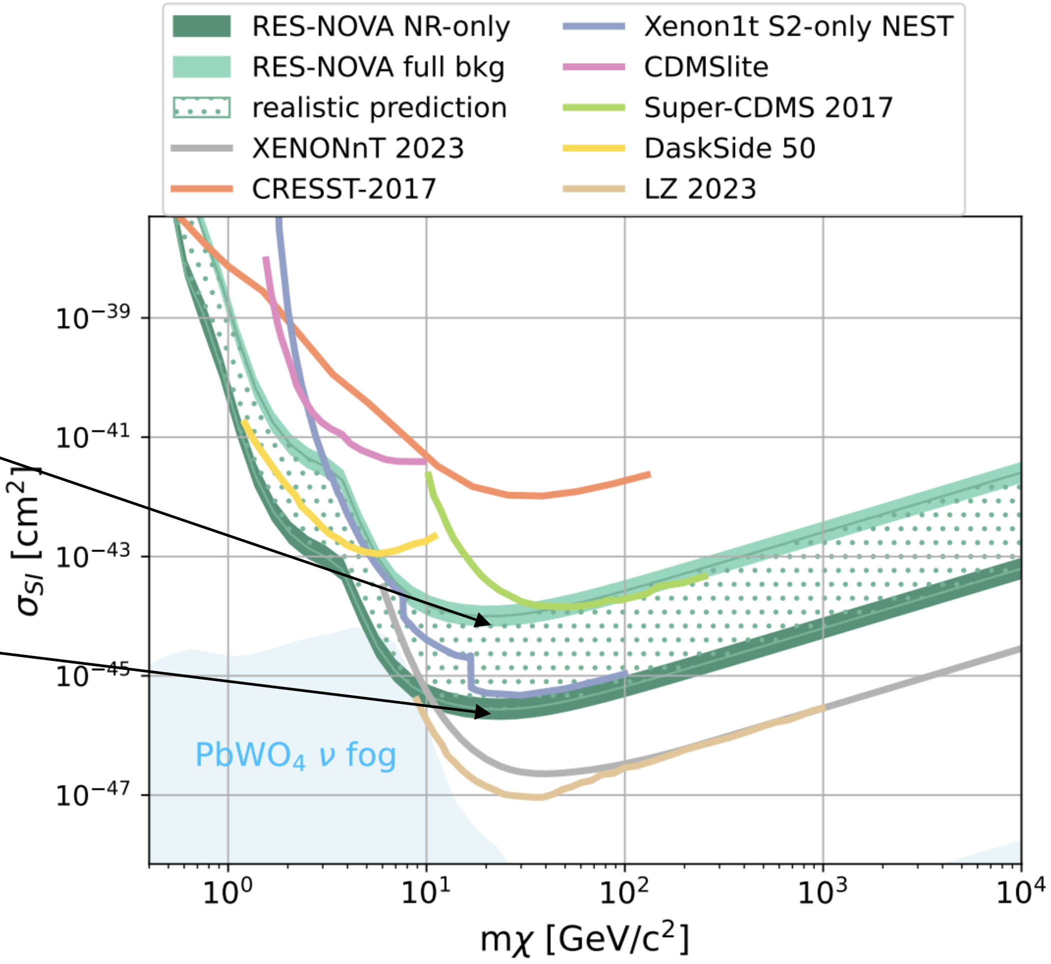




Projected sensitivity – spin independent

No bck rejection

Full e/gamma
bck rejection



RES-NOVA Coll., Phys. Rev. D 111, 103050 (2025)

Spin-dependent sensitivity

- ^{207}Pb has a natural abundance of 22.1% and it's a 2p1/2 neutron hole within the doubly magic ^{208}Pb

$$\mathcal{O}_1 = 1_\chi 1_N \qquad \mathcal{O}_9 = i\vec{S}_\chi \cdot \left(\vec{S}_N \times \frac{\vec{q}}{m_N}\right)$$

$$\mathcal{O}_3 = i\vec{S}_N \cdot \left(\frac{\vec{q}}{m_N} \times \vec{v}^\perp\right) \qquad \mathcal{O}_{10} = i\vec{S}_N \cdot \left(\frac{\vec{q}}{m_N}\right)$$

$$\mathcal{O}_4 = \vec{S}_\chi \cdot \vec{S}_N \qquad \mathcal{O}_{11} = i\vec{S}_\chi \cdot \left(\frac{\vec{q}}{m_N}\right)$$

$$\mathcal{O}_5 = i\vec{S}_\chi \cdot \left(\frac{\vec{q}}{m_N} \times \vec{v}^\perp\right) \qquad \mathcal{O}_{12} = \vec{S}_\chi \cdot (\vec{S}_N \times \vec{v}^\perp)$$

$$\mathcal{O}_6 = \left(\vec{S}_\chi \cdot \frac{\vec{q}}{m_N}\right)\left(\vec{S}_N \cdot \frac{\vec{q}}{m_N}\right) \qquad \mathcal{O}_{13} = i(\vec{S}_\chi \cdot \vec{v}^\perp)\left(\vec{S}_N \cdot \frac{\vec{q}}{m_N}\right)$$

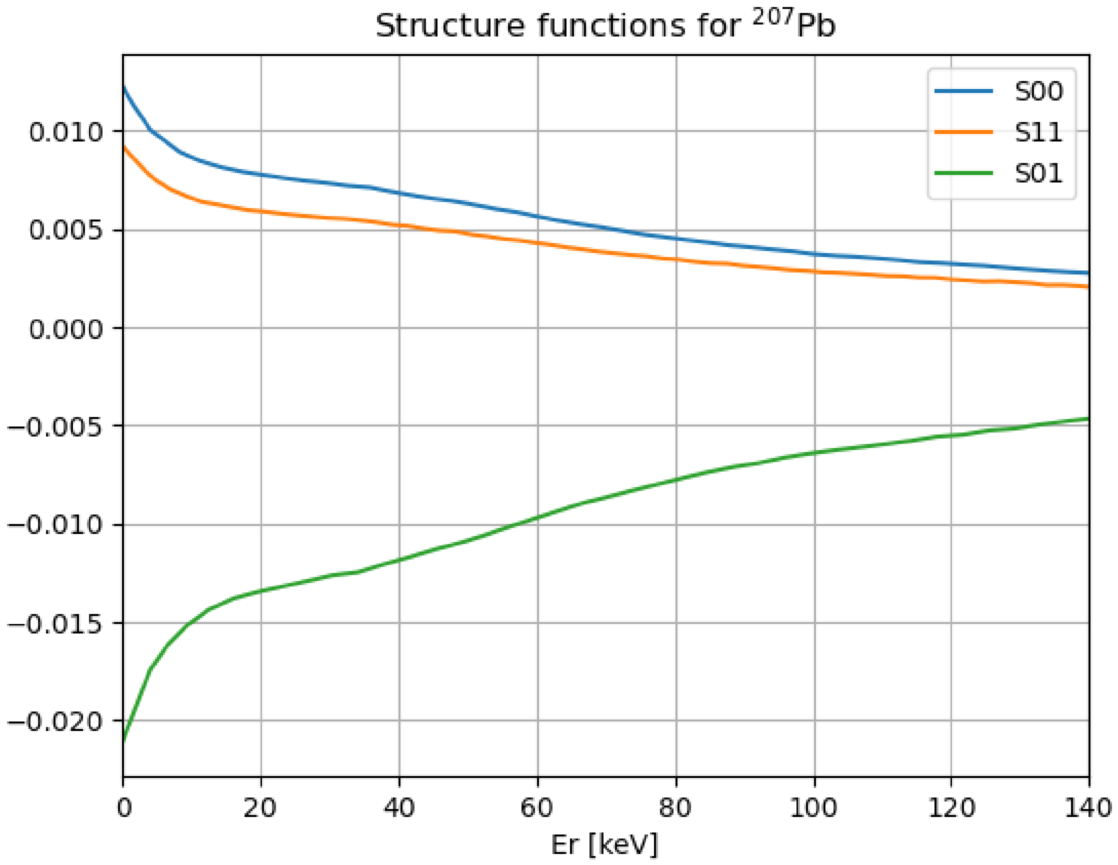
$$\mathcal{O}_7 = \vec{S}_N \cdot \vec{v}^\perp \qquad \mathcal{O}_{14} = i\left(\vec{S}_\chi \cdot \frac{\vec{q}}{m_N}\right)(\vec{S}_N \cdot \vec{v}^\perp)$$

$$\mathcal{O}_8 = \vec{S}_\chi \cdot \vec{v}^\perp \qquad \mathcal{O}_{15} = -\left(\vec{S}_\chi \cdot \frac{\vec{q}}{m_N}\right)\left[(\vec{S}_N \times \vec{v}^\perp) \cdot \frac{\vec{q}}{m_N}\right].$$

Spin-dependent sensitivity

- ^{207}Pb has a natural abundance of 22.1% and it's a 2p1/2 neutron hole within the doubly magic ^{208}Pb
- It's heavy and has known SD structure function => good target for heavy DM

- ^{17}O has 0.037% of abundance
- It's light and no known SD structure fucntion => kinematically favored target for light DM



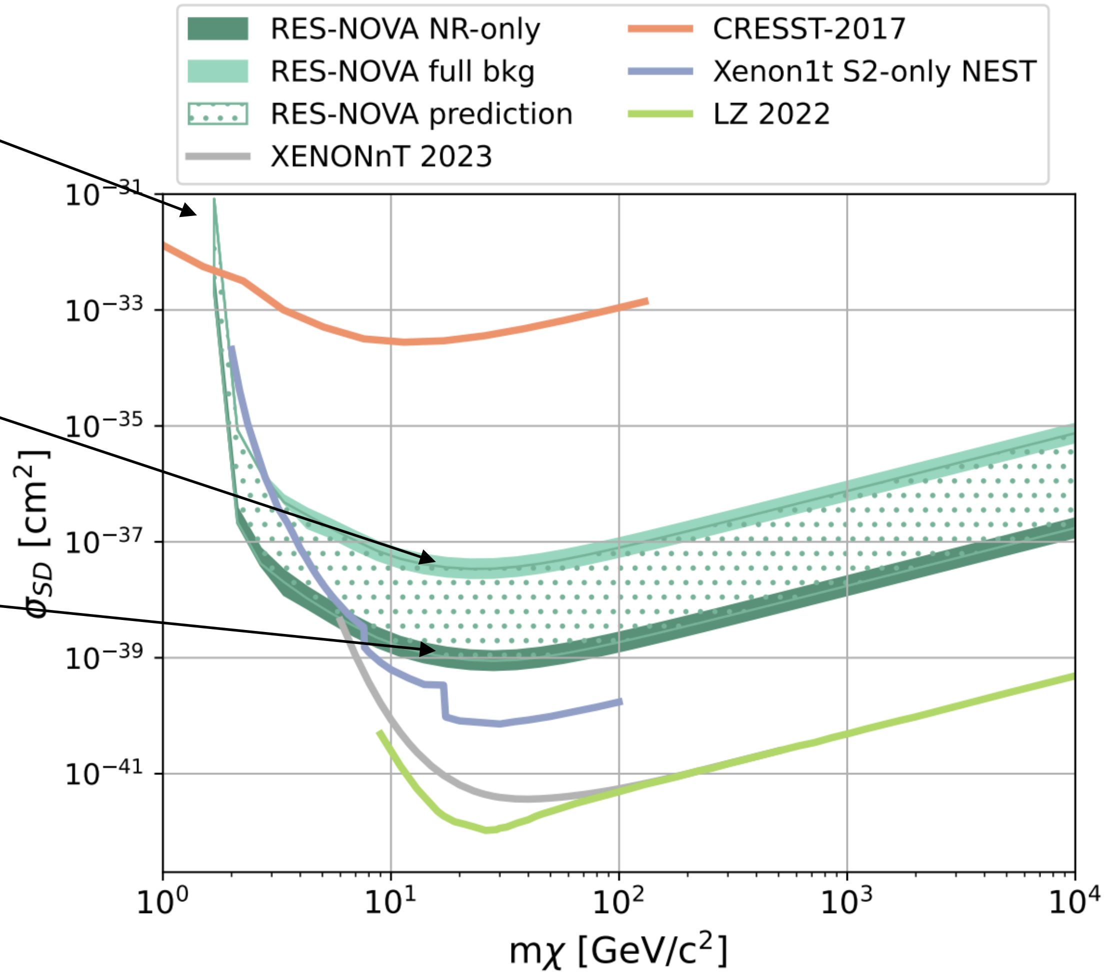
Kosmas, T. S. and Vergados, J. D. Cold dark matter in SUSY theories, 10.1103/PhysRevD.55.1752

Spin-dependent sensitivity

^{17}O leading sensitivity

No bck rejection

Full e/gamma bck rejection



RES-NOVA Coll., Phys. Rev. D 111, 103050 (2025)

Why we can do SN and DM without spectroscopy

Sei V das gesamte Volumen des Gases, T die mittlere lebendige Kraft eines Gasmoleküls und N die Gesamtzahl aller Moleküle des Gases, endlich m die Masse eines Gasmoleküls, so ist für den Zustand des Wärmegleichgewichtes

$$f(x, y, z, u, v, w) = \frac{N}{V\left(\frac{4\pi T}{3m}\right)^{\frac{3}{2}}} \cdot e^{-\frac{3m}{4T}(u^2+v^2+w^2)}$$

Substituiert man diesen Wert in Gleichung (61), so erhält man

(62) $$\Omega = \frac{3N}{2} + Nl\left[V\left(\frac{4\pi T}{3m}\right)^{\frac{3}{2}}\right] - NlN.$$

Versteht man nun unter dQ das dem Gase zugeführte Wärmedifferentiale, so ist

(63) $$dQ = NdT + pdV$$

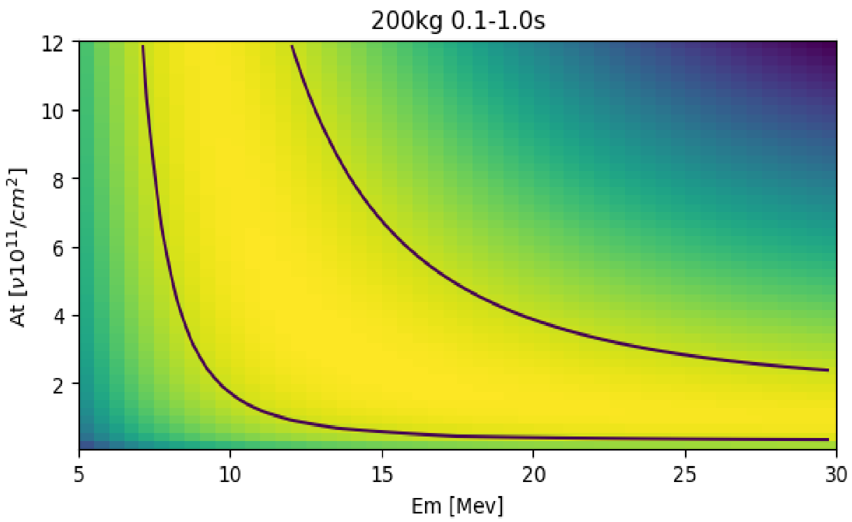
und

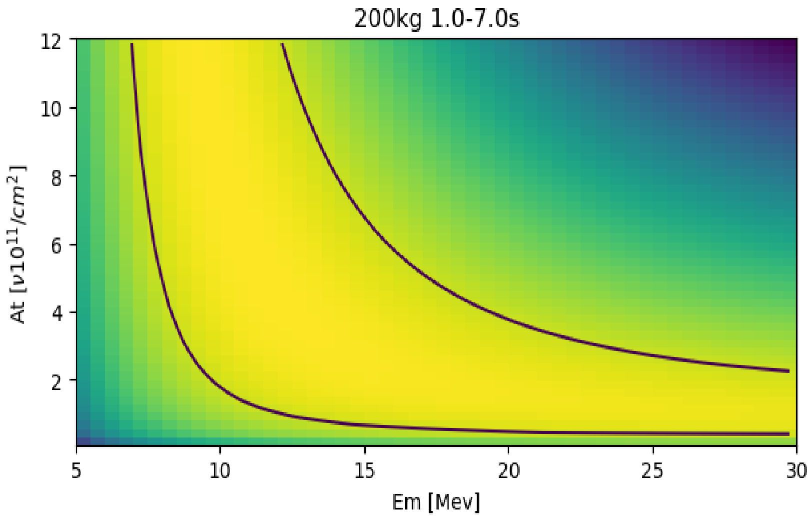
(64) $$pV = \frac{2N}{3} \cdot T.$$

p ist der Druck, bezogen auf die Flächeneinheit. Die Entropie des Gases ist dann:

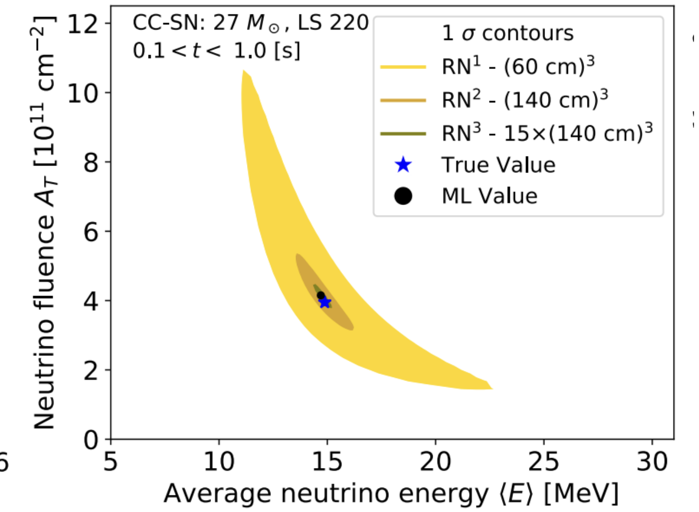
$$\int \frac{dQ}{T} = \frac{2}{3} N \cdot l(V \cdot T^{\frac{3}{2}}) + C.$$

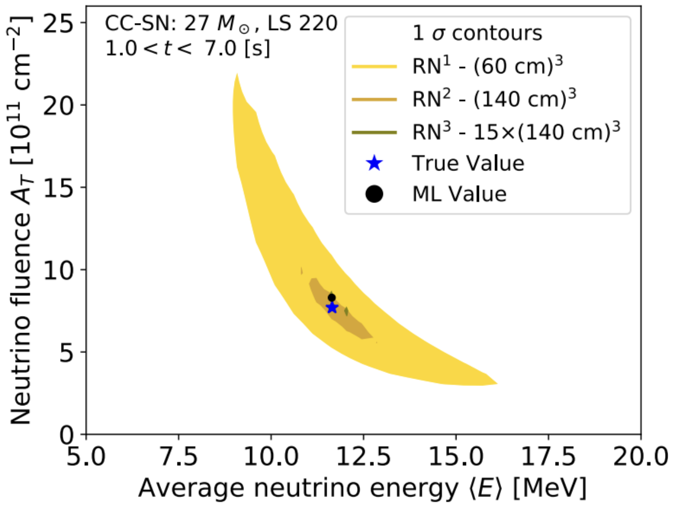
Constraining SN parameters





~200kg coming in the next 3-5 years





RN^1 2.4t

RN^2 31t

RN^3 465t

Low background observatory featuring:

- ^{17}O → Light, spin-dependent DM (q→0), Solar ν CEνES
- ^{207}Pb → Spin-dependent DM, known spin structure, CEνES
- ^{208}Pb → Scalar DM, CEνES
- ^{180}W → Scalar DM, CEνES, double EC
- ^{186}W → Scalar DM, CEνES, double beta decay

Conclusions

- The high Z (N) of Pb makes it ideal for DM direct search
- The isotopic abundance of ^{207}Pb enables comprehensive exploration of DM-nucleus couplings
- Archaeo-Pb is a great material for low background observatories
- The various isotopes present in $PbWO_4$ offer several interesting possibilities in several fields of (astro) particle physics

Thank you!