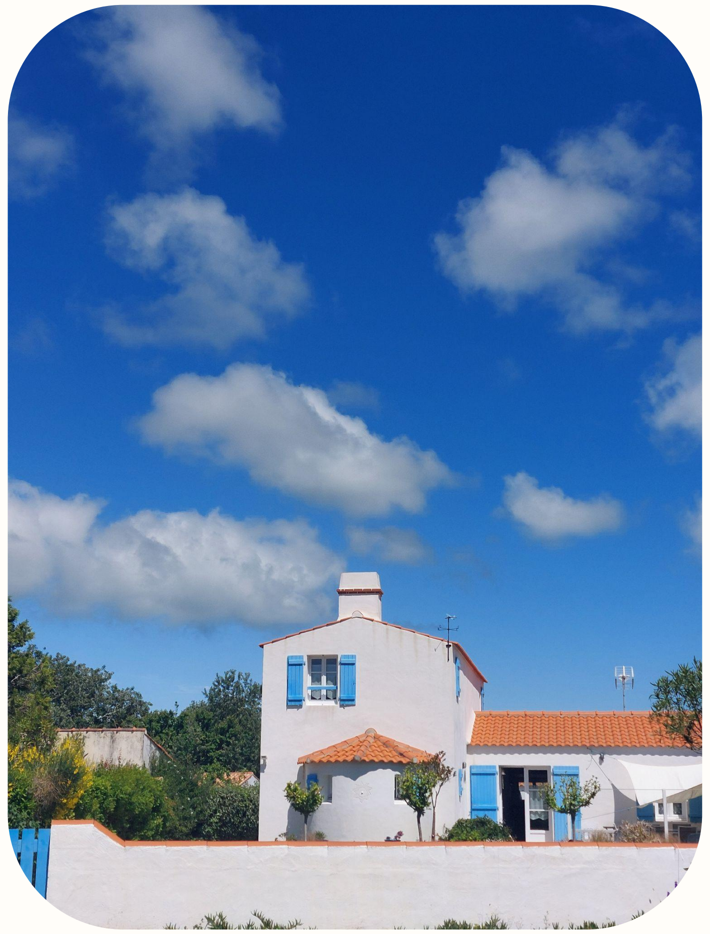
EXPLORING POTENTIAL:
ALPS II'S TES DETECTION SYSTEM FOR DIRECT DARK MATTER INVESTIGATIONS



Christina Schwemmbauer[1],

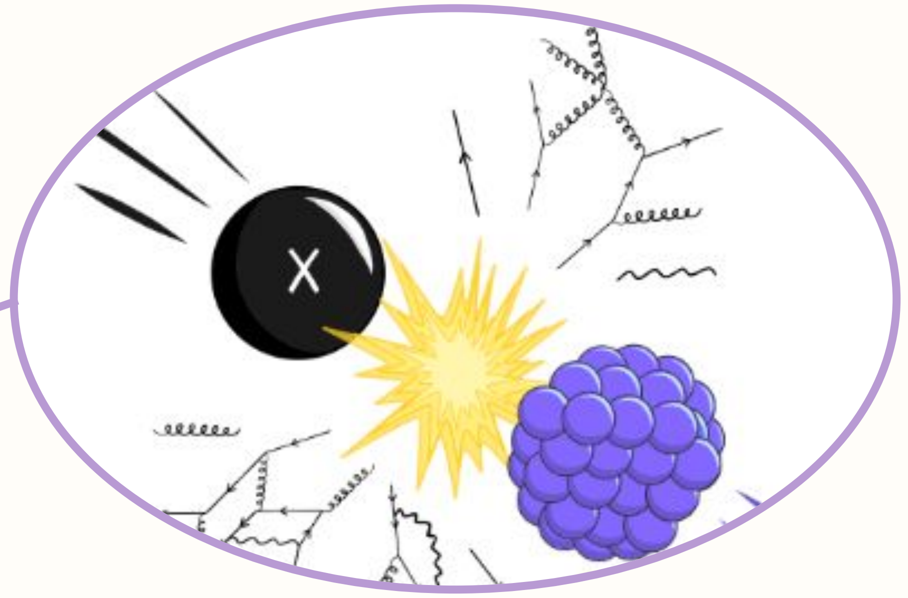
Katharina-Sophie Isleif[1, now at 2], Friederike Januschek[1], Axel Lindner[1], Manuel Meyer[3], Gulden Othman[4], Elmeri Rivasto[3], José Alejandro Rubiera Gimeno[1]

[1]Deutsches Elektronen Synchrotron DESY, Hamburg, Germany
[2]Helmut-Schmidt Universität, Hamburg, Germany
[3]Southern Denmark University, Odense, Denmark
[4]Institut für Experimentalphysik, Universität Hamburg, Germany

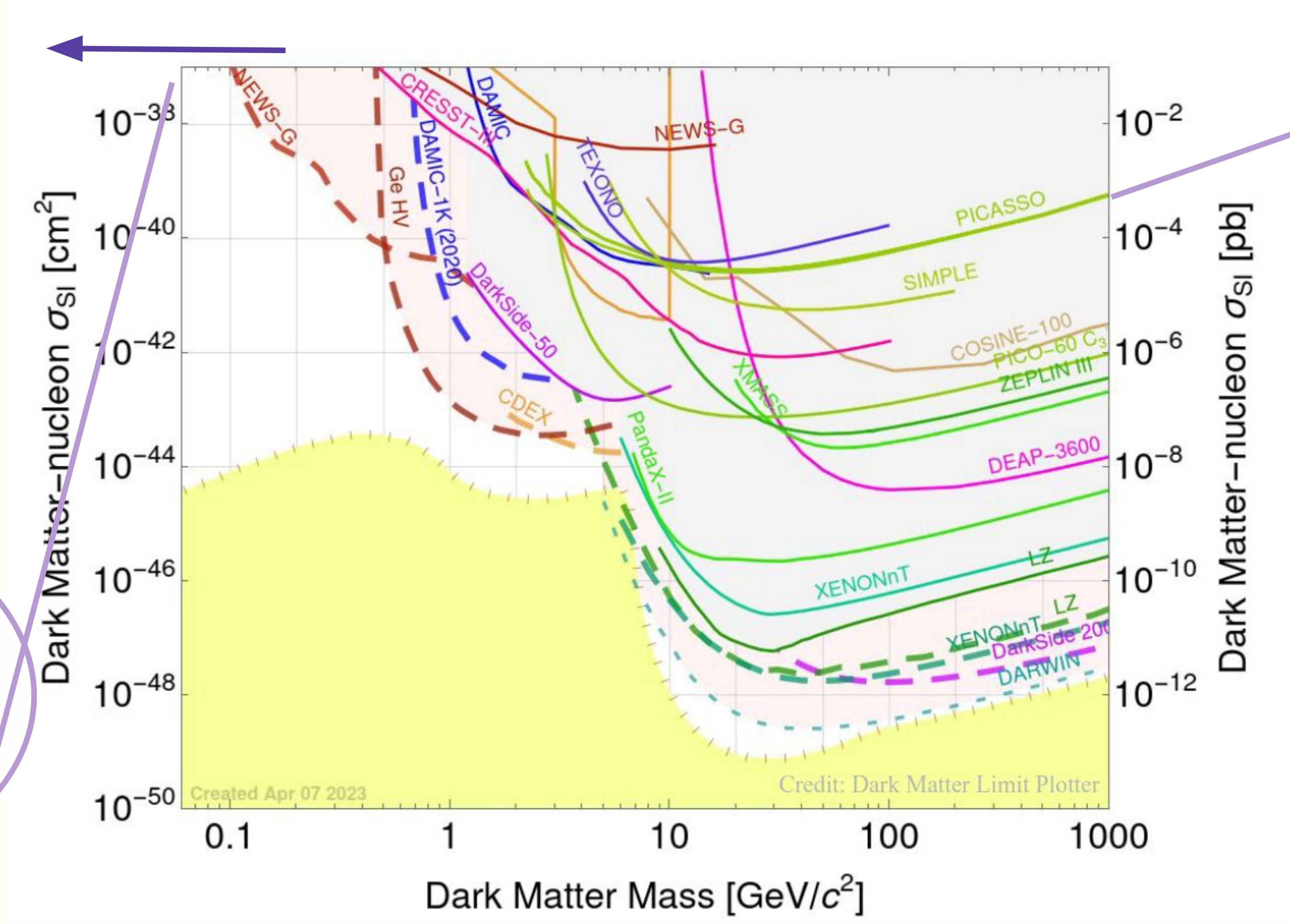
EDSU-Tools 2024

Direct Detection of sub-GeV DM?

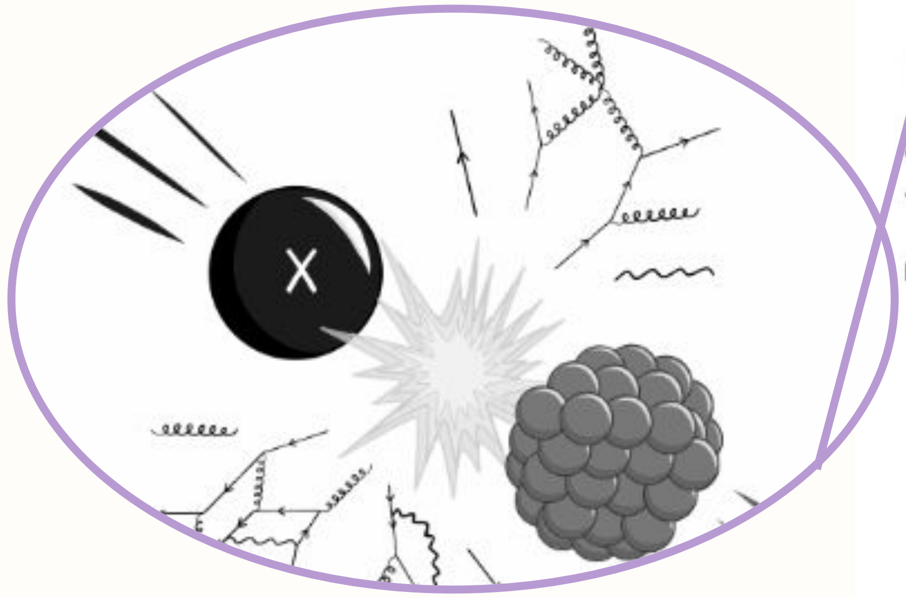
Sketch adapted from Benjamin V. Lehmann



DM-nucleon scattering



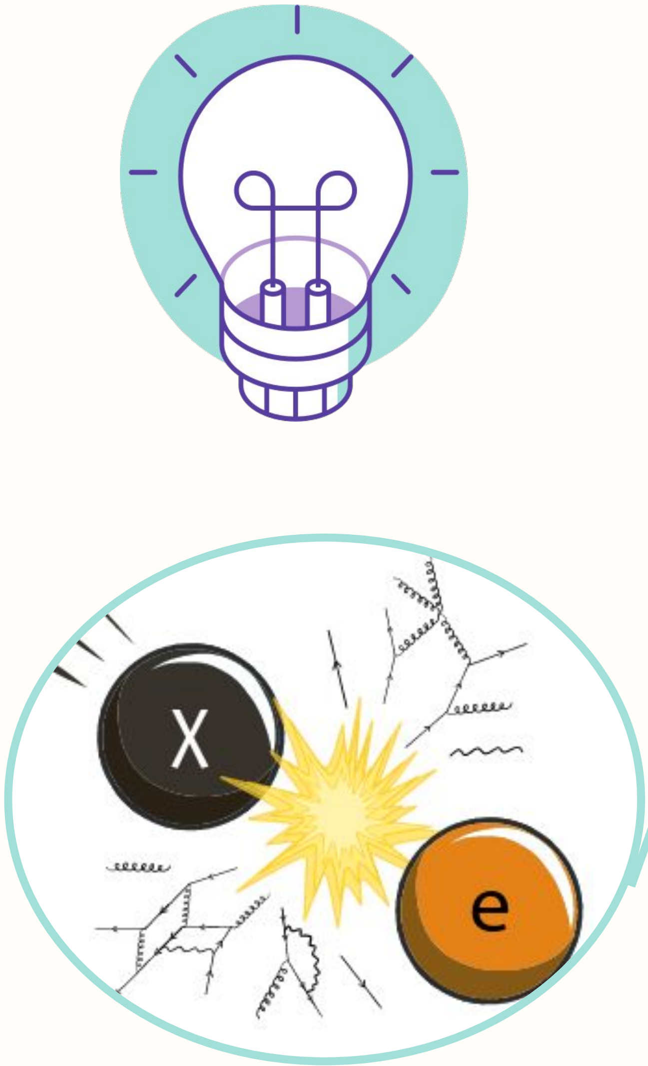
from S. Lindemann, “WIMP direct detection experiments”, 18th PATRAS Workshop on Axions, WIMPs and WISPs, Rijeka 2023



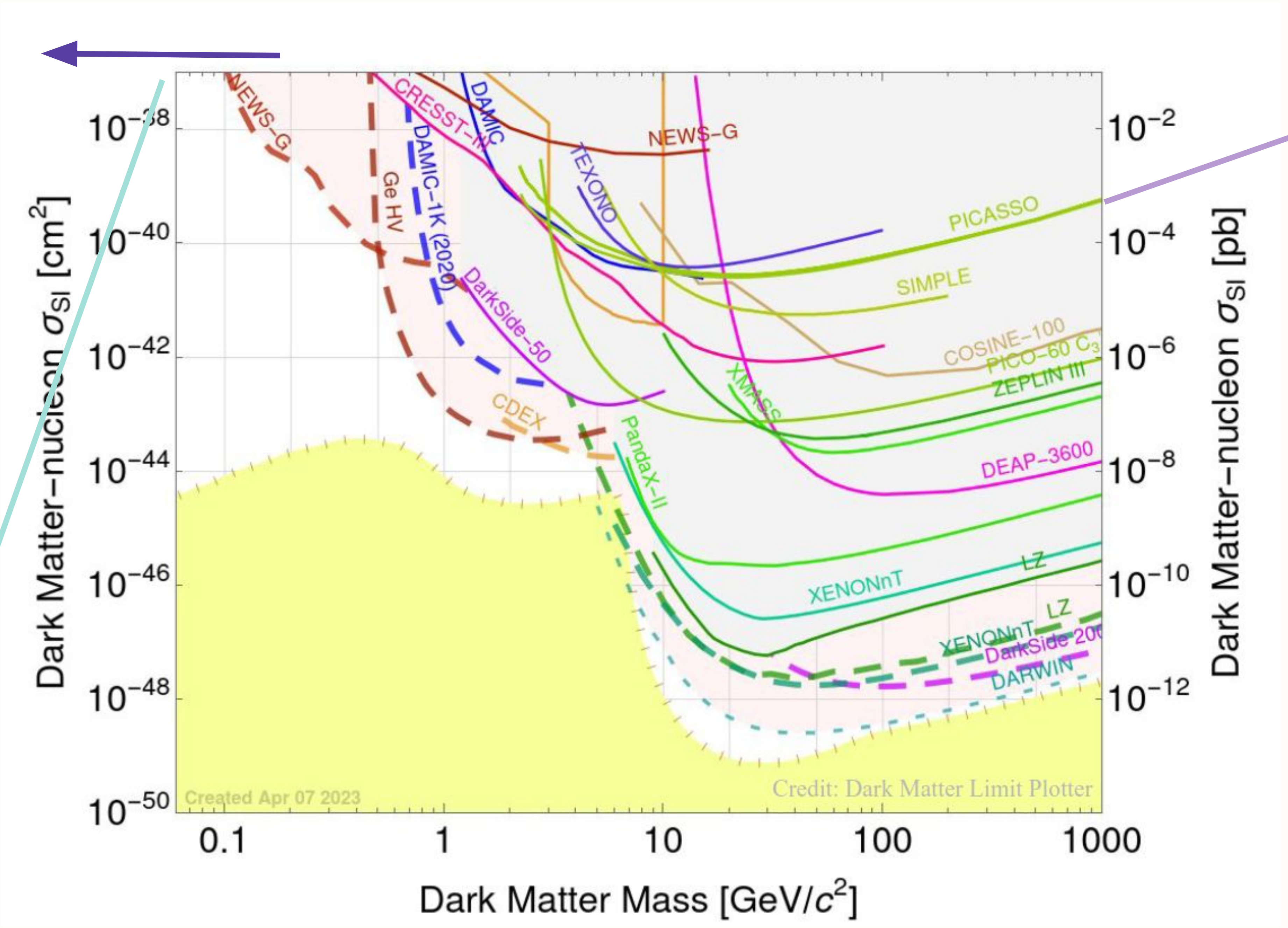
very low recoil energy
→ experiments lose sensitivity

Direct Detection of sub-GeV DM?

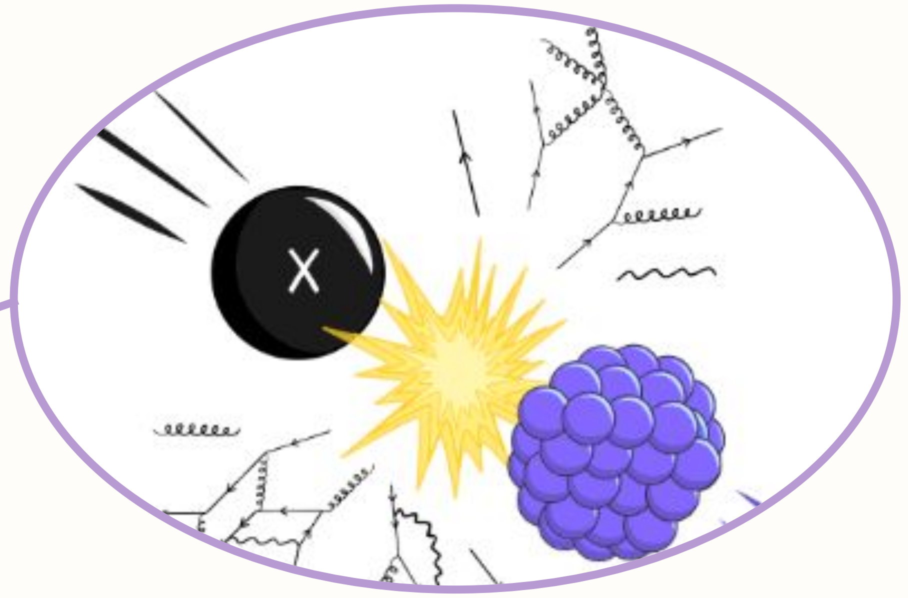
Sketch adapted from Benjamin V. Lehmann



DM-electron scattering

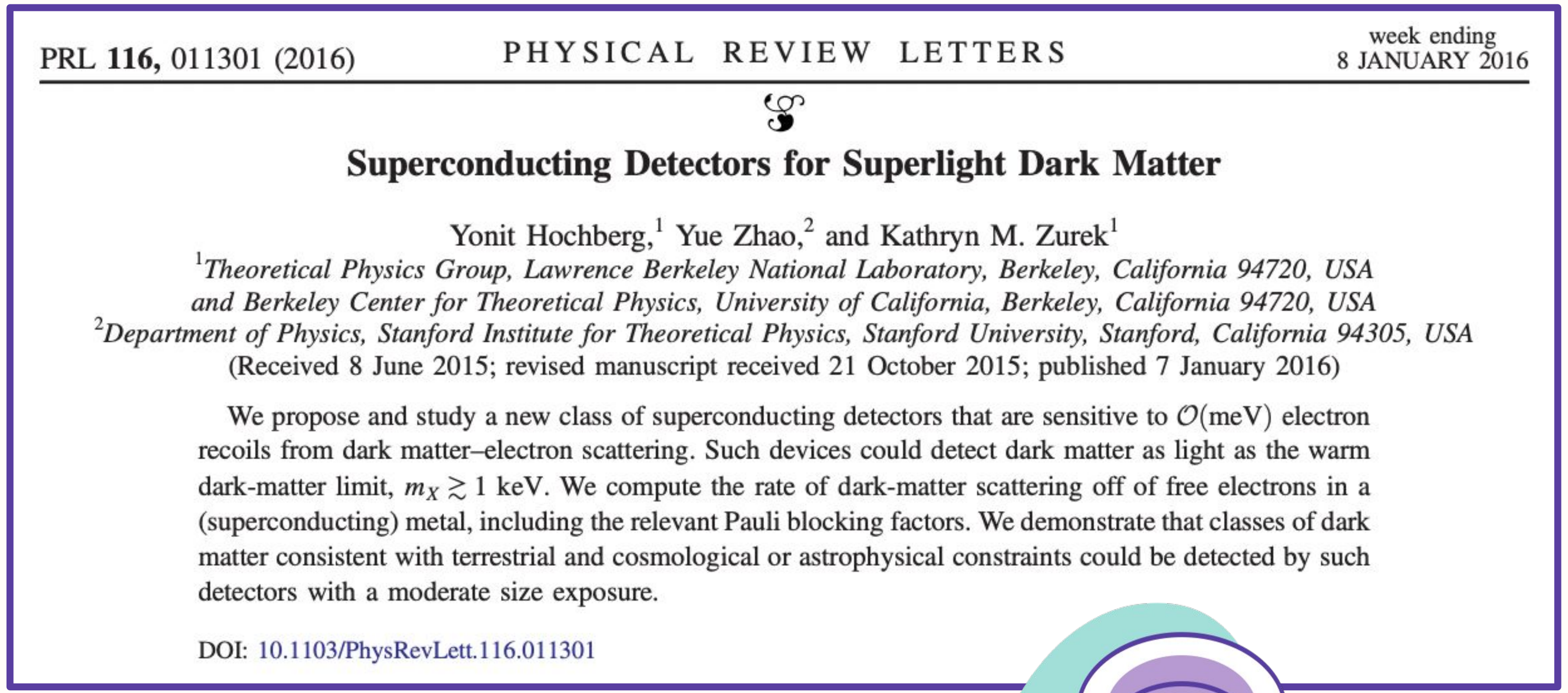


from S. Lindemann, "WIMP direct detection experiments", 18th PATRAS Workshop on Axions, WIMPs and WISPs, Rijeka 2023



DM-nucleon scattering

Superconducting Detectors



PRL **116,** 011301 (2016) PHYSICAL REVIEW LETTERS week ending 8 JANUARY 2016

Superconducting Detectors for Superlight Dark Matter

Yonit Hochberg,[1] Yue Zhao,[2] and Kathryn M. Zurek[1]

[1]*Theoretical Physics Group, Lawrence Berkeley National Laboratory, Berkeley, California 94720, USA and Berkeley Center for Theoretical Physics, University of California, Berkeley, California 94720, USA*

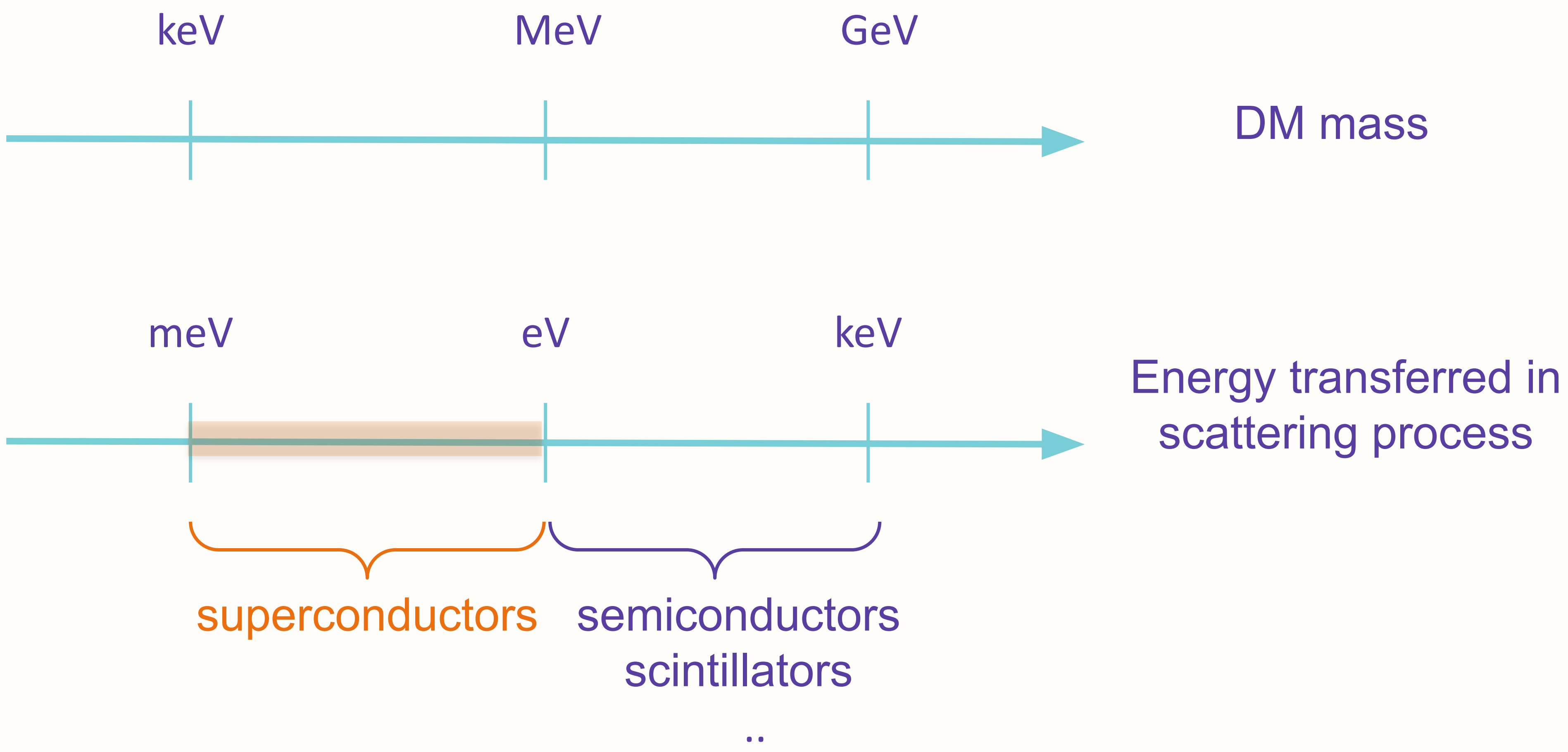
[2]*Department of Physics, Stanford Institute for Theoretical Physics, Stanford University, Stanford, California 94305, USA*

(Received 8 June 2015; revised manuscript received 21 October 2015; published 7 January 2016)

We propose and study a new class of superconducting detectors that are sensitive to $\mathcal{O}$(meV) electron recoils from dark matter–electron scattering. Such devices could detect dark matter as light as the warm dark-matter limit, $m_X \gtrsim 1$ keV. We compute the rate of dark-matter scattering off of free electrons in a (superconducting) metal, including the relevant Pauli blocking factors. We demonstrate that classes of dark matter consistent with terrestrial and cosmological or astrophysical constraints could be detected by such detectors with a moderate size exposure.

DOI: 10.1103/PhysRevLett.116.011301

$$E_{\mathrm{T_{max}}} = E_{\mathrm{kin}} \sim m_\chi v^2 \sim 10^{-6} m_\chi$$



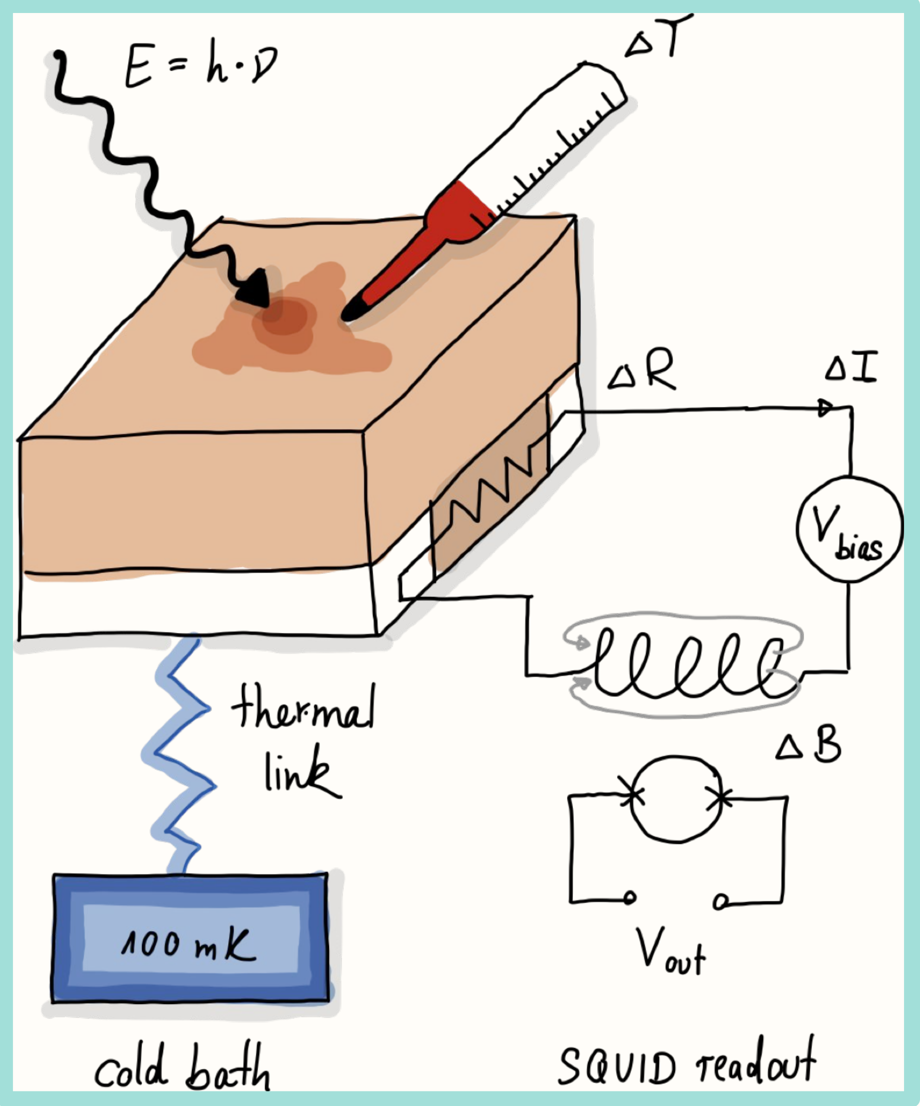
Transition Edge Sensors (TES)

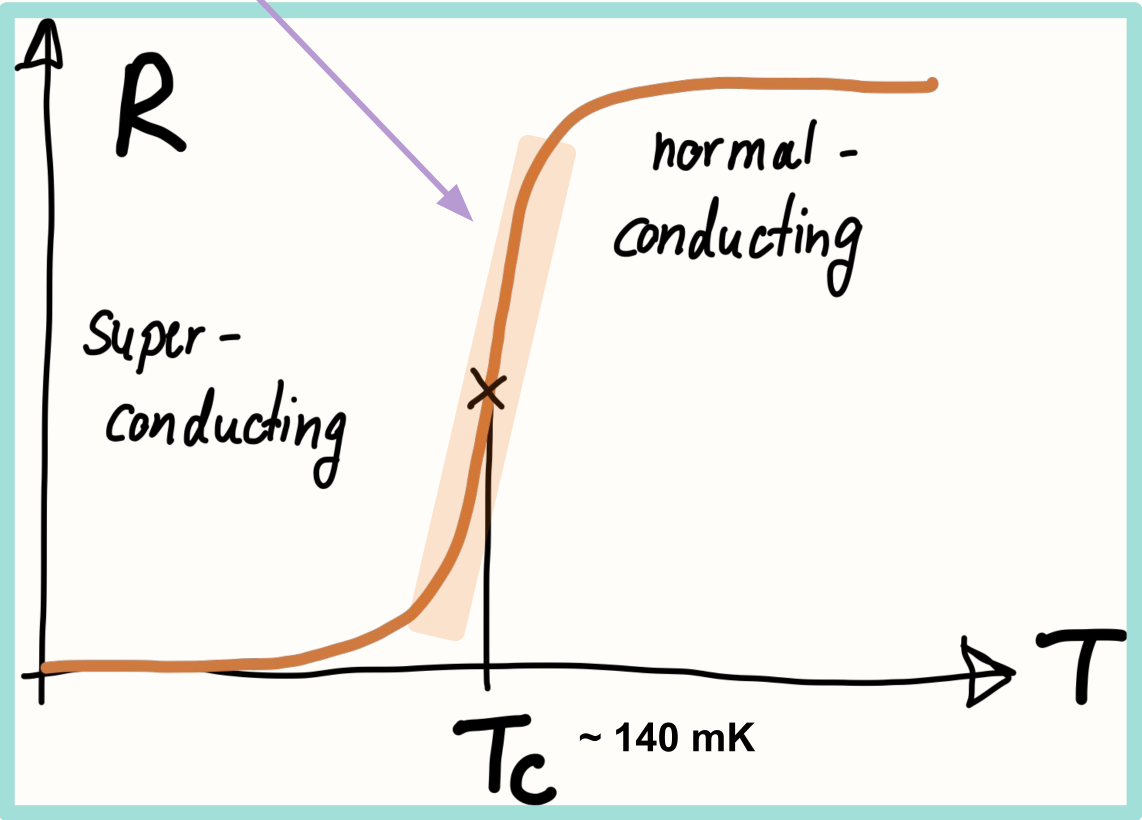
Incident photon on superconducting material leads to temperature increase

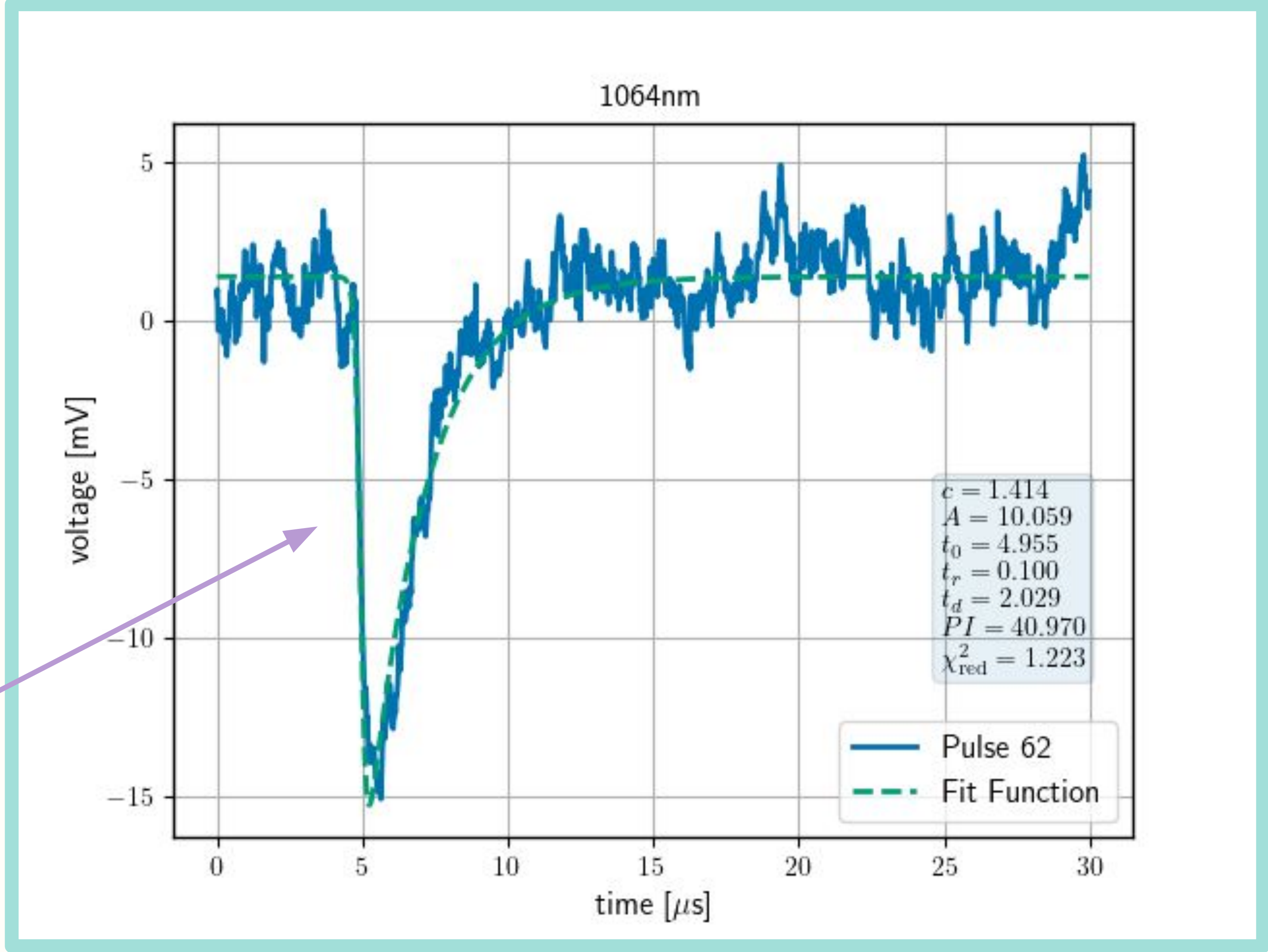
Leading to a large increase in resistance on transition curve

Single-photon pulse integral **proportional** to photon energy (1064 nm ≙ 1.165 eV)

Drawings courtesy of Katharina-Sophie Isleif



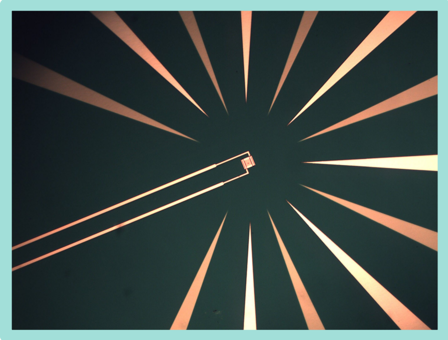


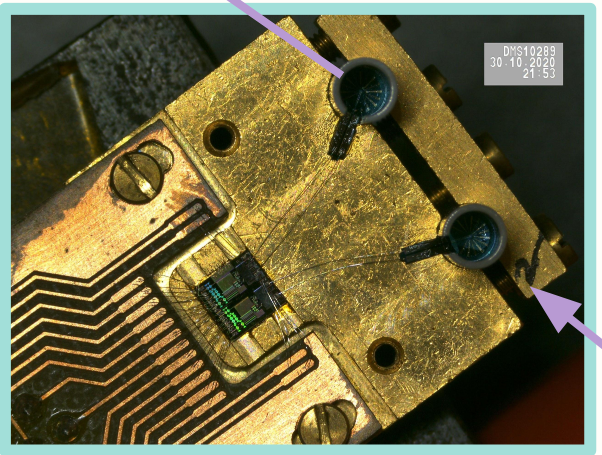


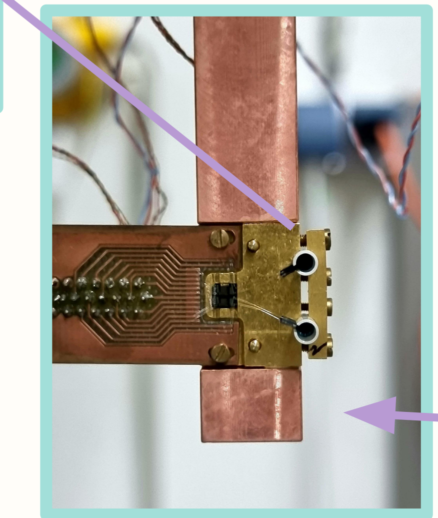
SQUID ≙ Superconducting Quantum Interference Device

TES Setup @ DESY.

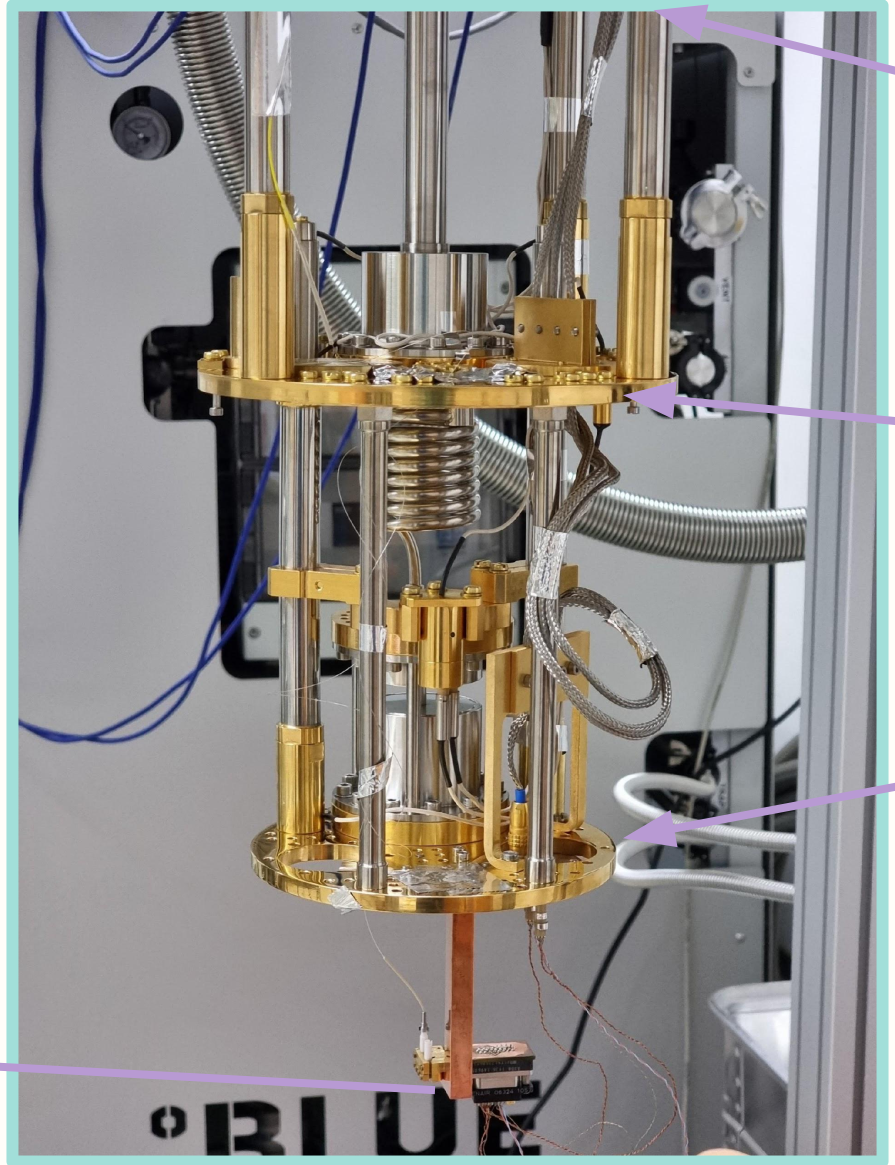
Area = 25 µm x 25 µm
Thickness = 20 nm

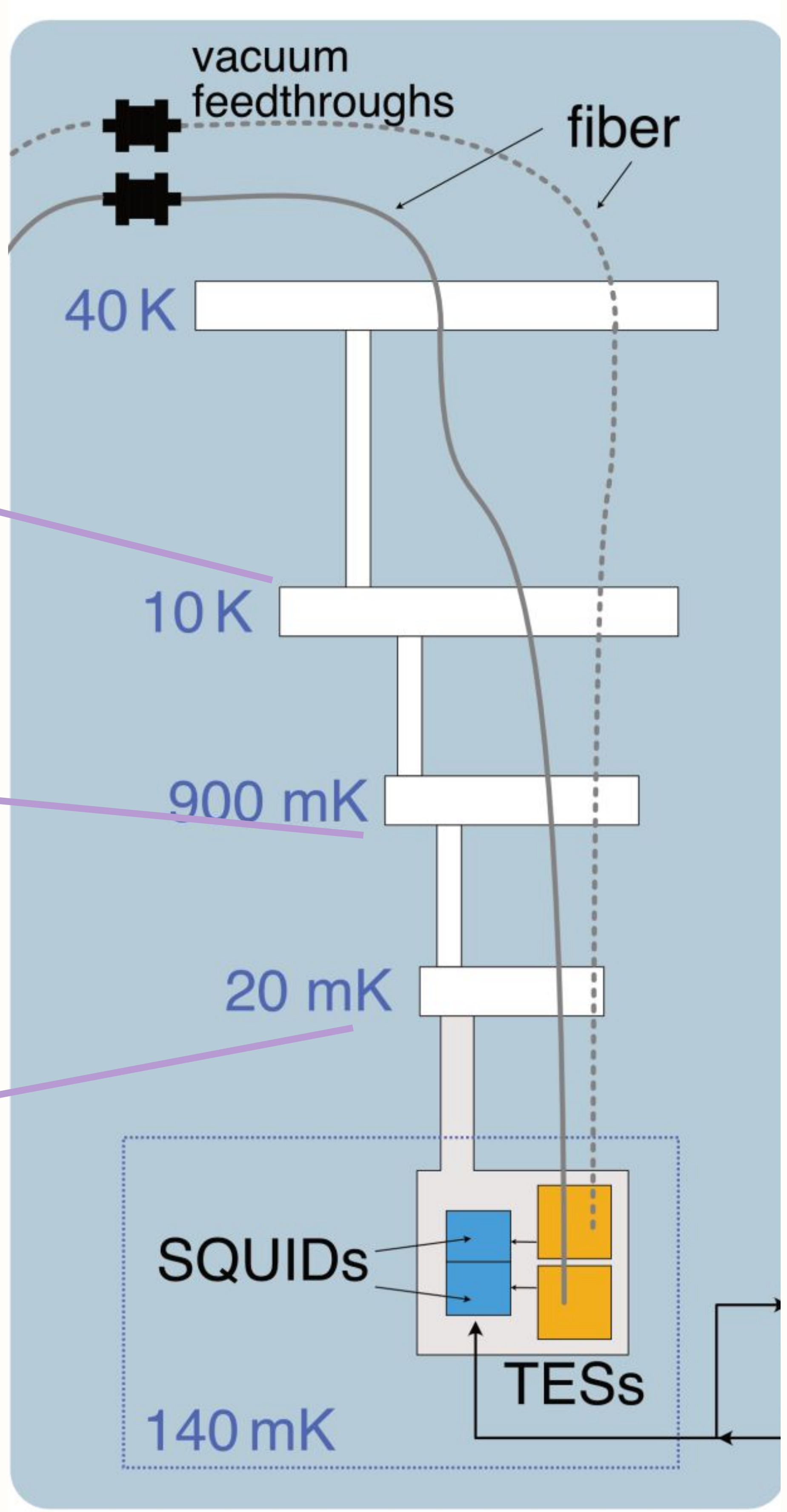






Inside of dilution refrigerator

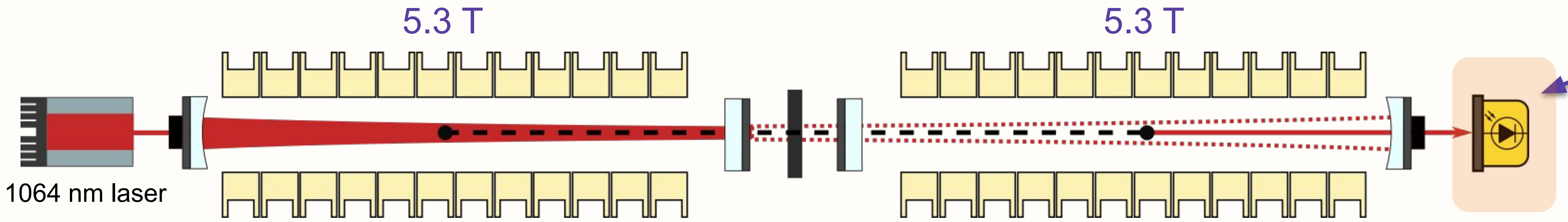




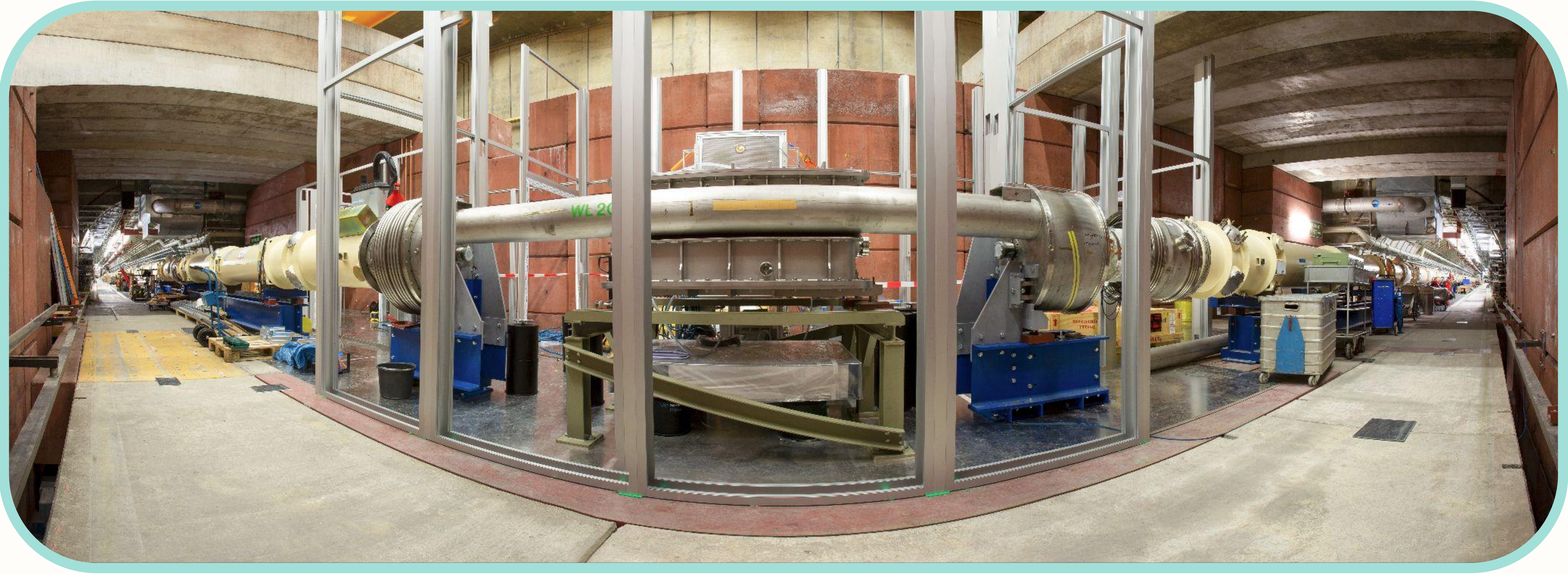
Sensors provided by: NIST

Packaging and SQUIDs provided by: PTB

TES at ALPS II



Light Shining through a Wall experiment
Challenge: Detect single photon from axion-photon conversion
Currently: HETerodyne Sensing
Future Option: TES



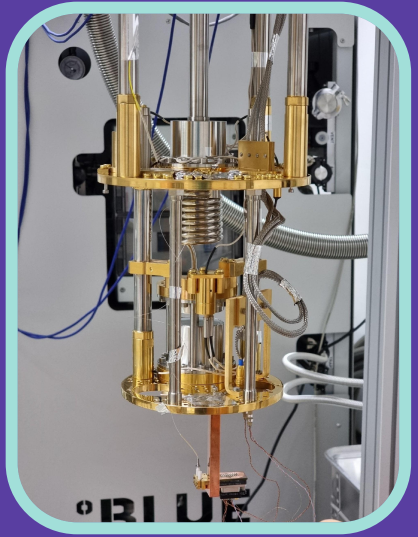
250m long experiment in “old” HERA accelerator tunnel

ALPS @ DESY.

See backup slides for details on ALPS II

ALPS II TES Requirements

Optimized setup, detector and analysis for 1064 nm (1.165 eV) photons

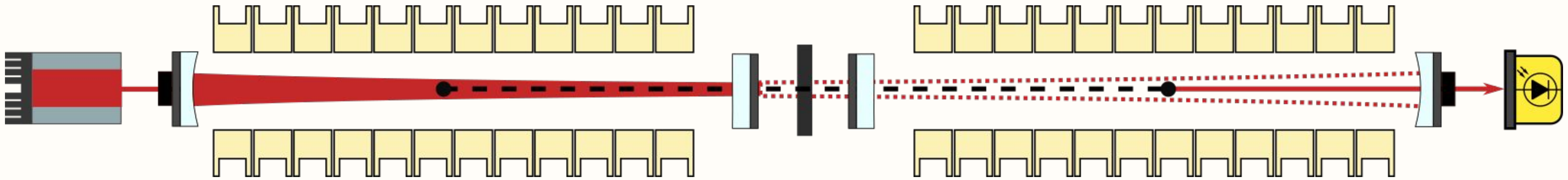


Extremely low background needed to detect ~ 1 photon/day

Currently 6.9x10^{-6} cps [1] intrinsic background

Very good system detection efficiency >50%

>90% measured in the lab! [2]

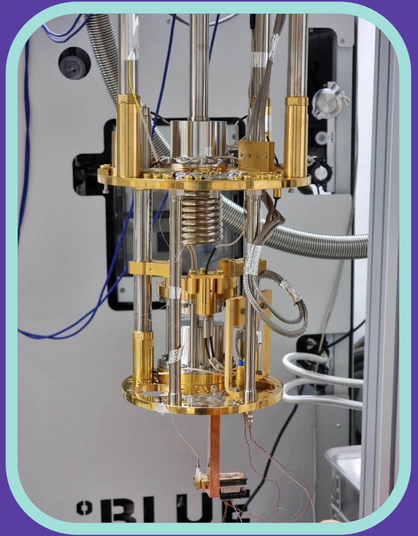


[1] R. Shah et al., PoS, EPS-HEP2021, 801 (2022)

[2] J.A. Rubiera Gimeno et al., PoS EPS-HEP2023, 567 (2024)

ALPS II DDM TES Requirements

Calibrated setup for a larger **energy range** especially **sub-eV**

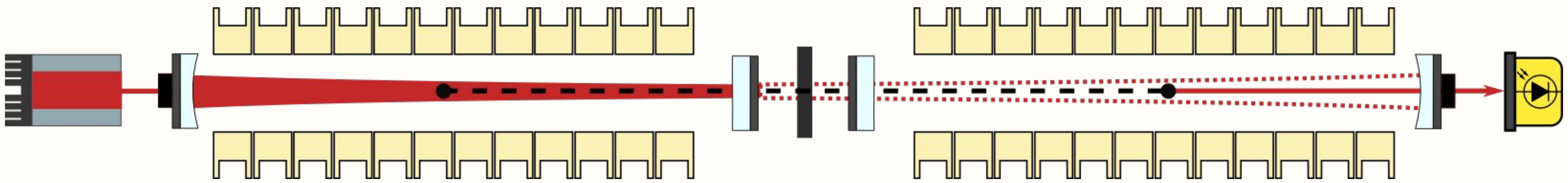


Extremely low background needed over a **range of energies**

Currently 6.9x10^{-6} cps [1] intrinsic background @ 1064 nm

Good energy resolution over broad energy spectrum

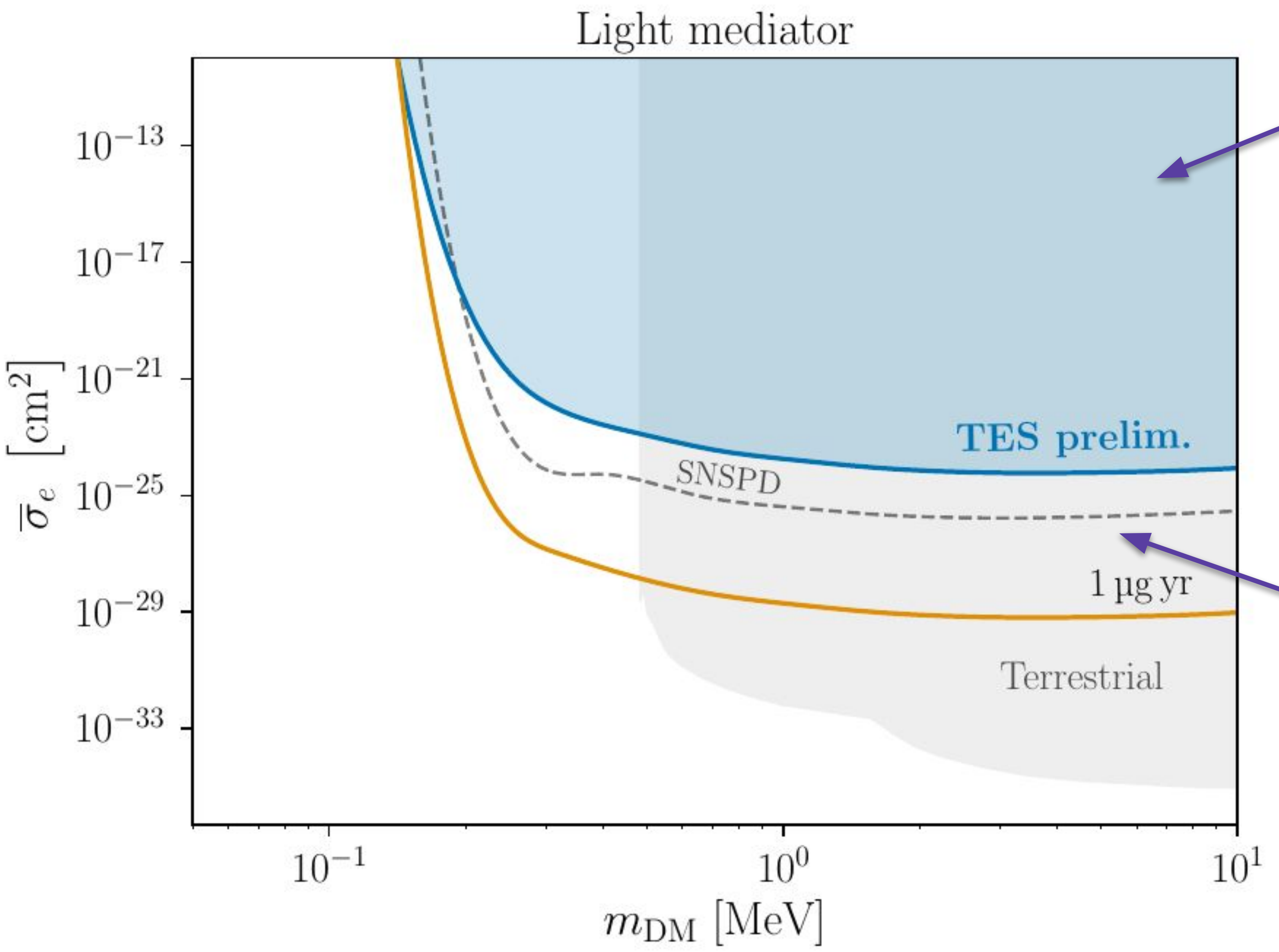
~5% @ 1.165 eV [2]



[1] R. Shah et al., PoS, EPS-HEP2021, 801 (2022)

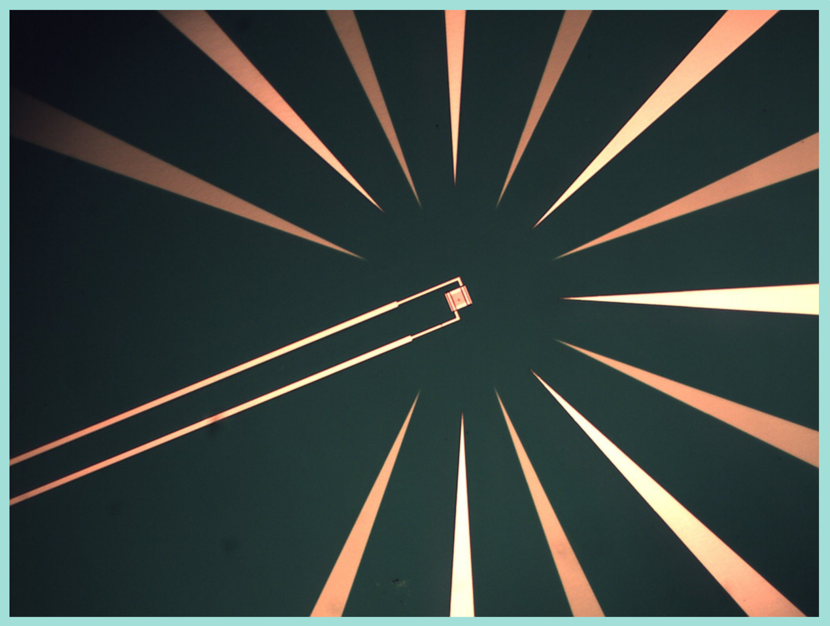
[2] J.A. Rubiera Gimeno et al., PoS EPS-HEP2023, 567 (2024)

TES for DDM



Possible sensitivity based on our TES setup (based on previous intrinsic measurements)
→ **DDM hits expected to look like photon signature**

Superconducting Nanowire Single Photon Detector (SNSPD) with 4 dark counts in 180 h with 0.73 eV energy threshold [3]



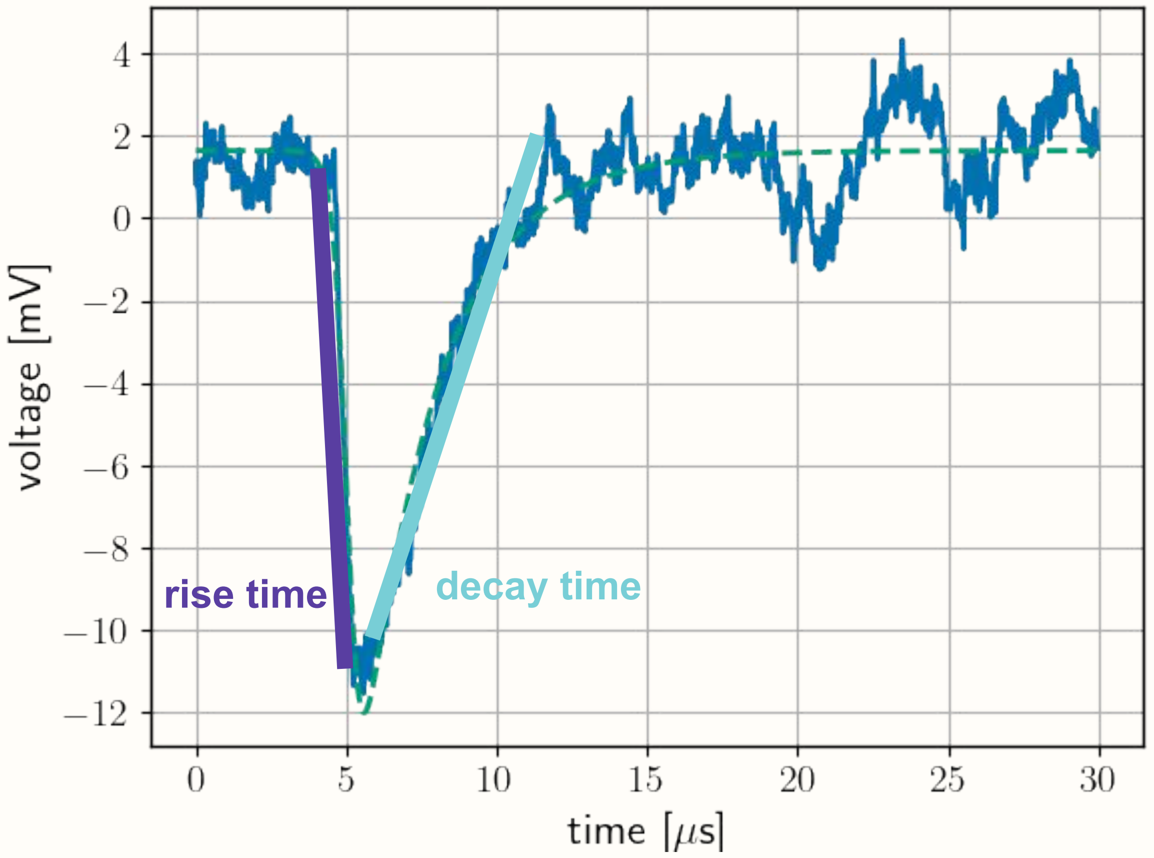
Challenges:
- low mass (0.2 ng)
- small area (25 x 25 µm)
- **limited knowledge about broadband response**

Projections and plot by Benjamin V. Lehmann

[3]Hochberg, Y. et al. arXiv:2110.01586 (2021)

Fitting procedure

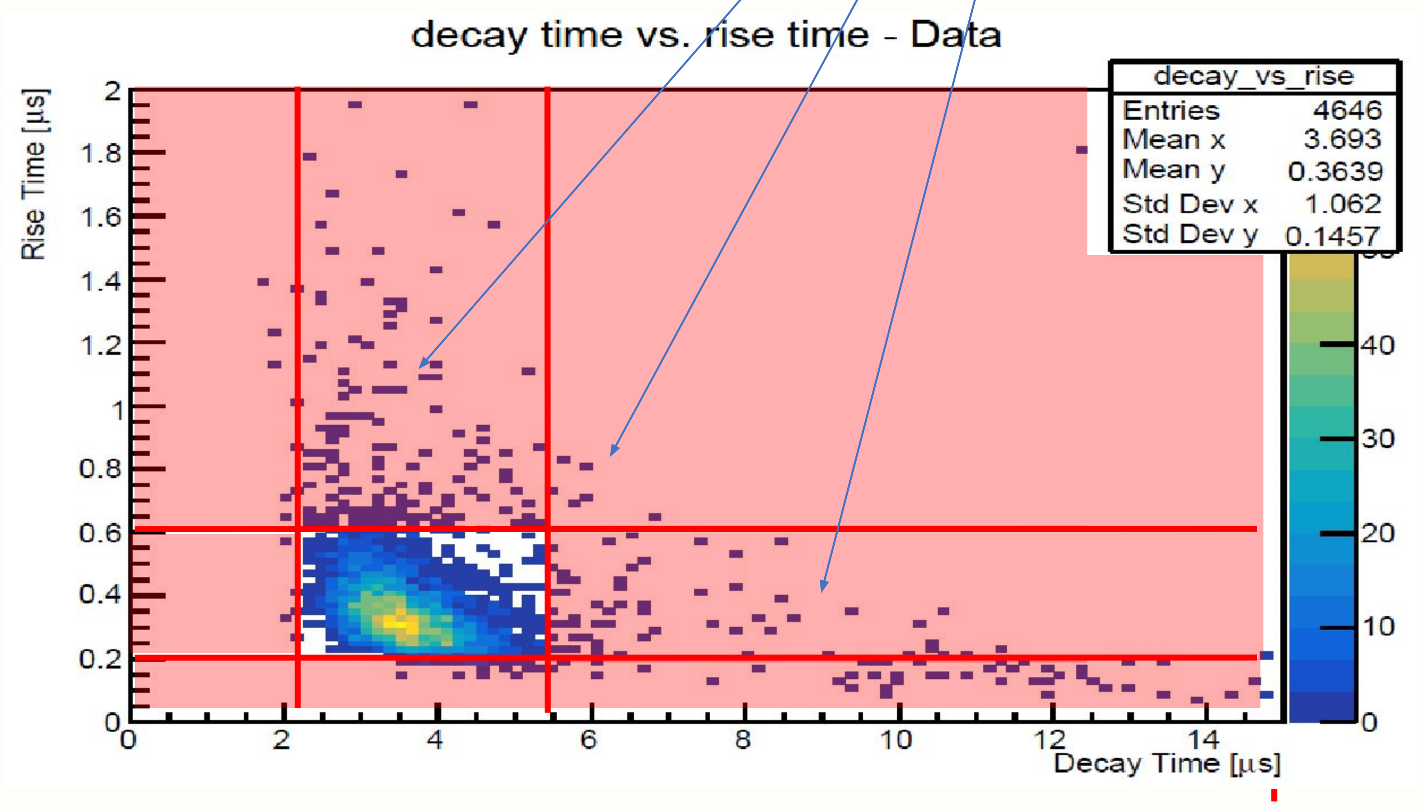
cutting away signals outside of central region

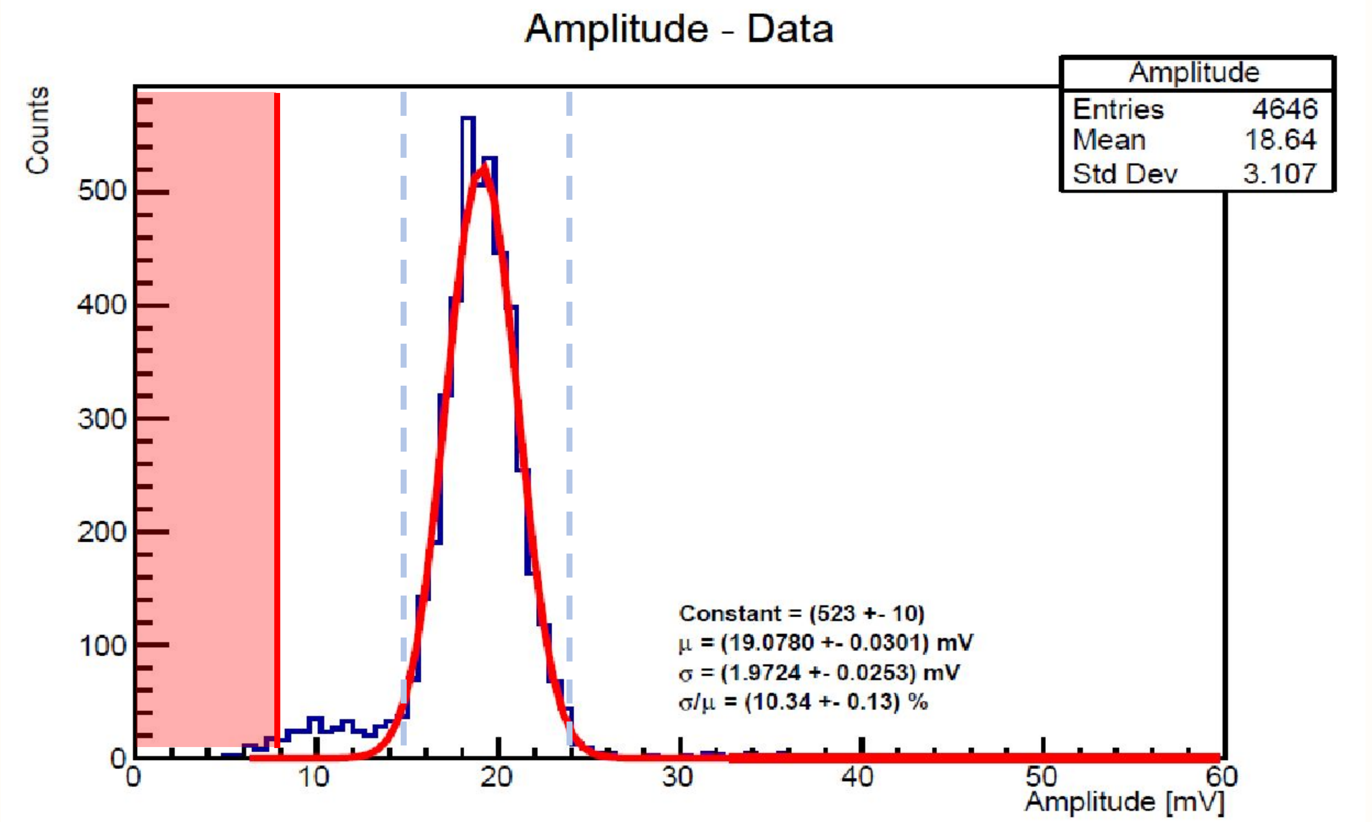


$$U(t) = -\frac{2A}{e^{-\frac{1}{\tau_{\text{rise}}}(t-t_0)} + e^{\frac{1}{\tau_{\text{decay}}}(t-t_0)}} + V_0$$

- 1064 nm calibration laser
- pulse shape fitted to all triggered pulses
- apply cuts on parameters like rise and decay time, pulse height or amplitude

→ apply for multiple wavelengths/energies?



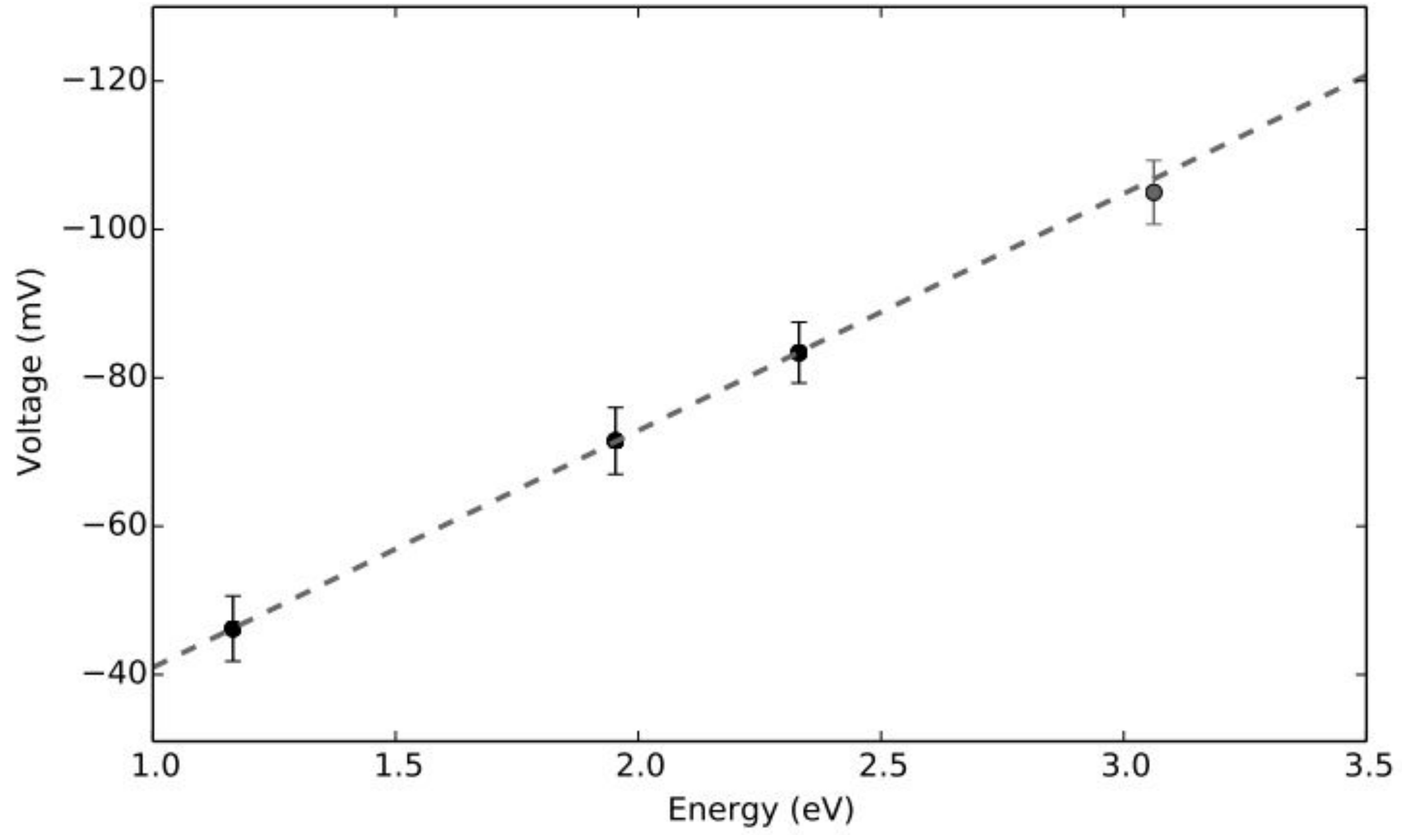


Energy Calibration

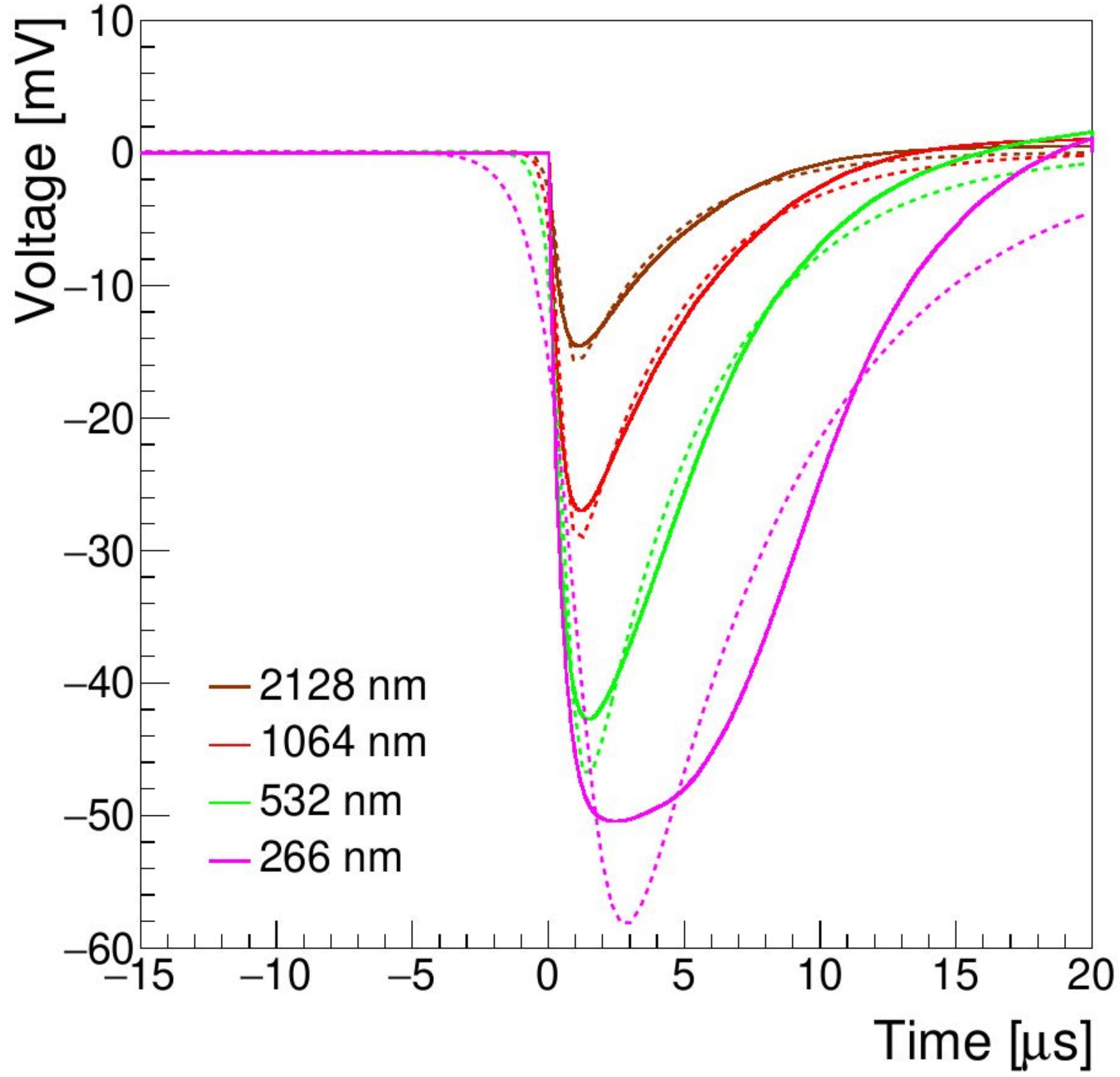
Expectations:
- linear relationship between pulse integral/amplitude and energy
- approximately constant behavior of rise and decay time (mainly governed by TES circuit)
- more challenging at lower energies (noise, etc.)

→previous simulations (DOI:10.22323/1.454.0055) showed promising results for a sub-MeV DM search based on these assumptions

Calibration measurement of a previous setup (2016)

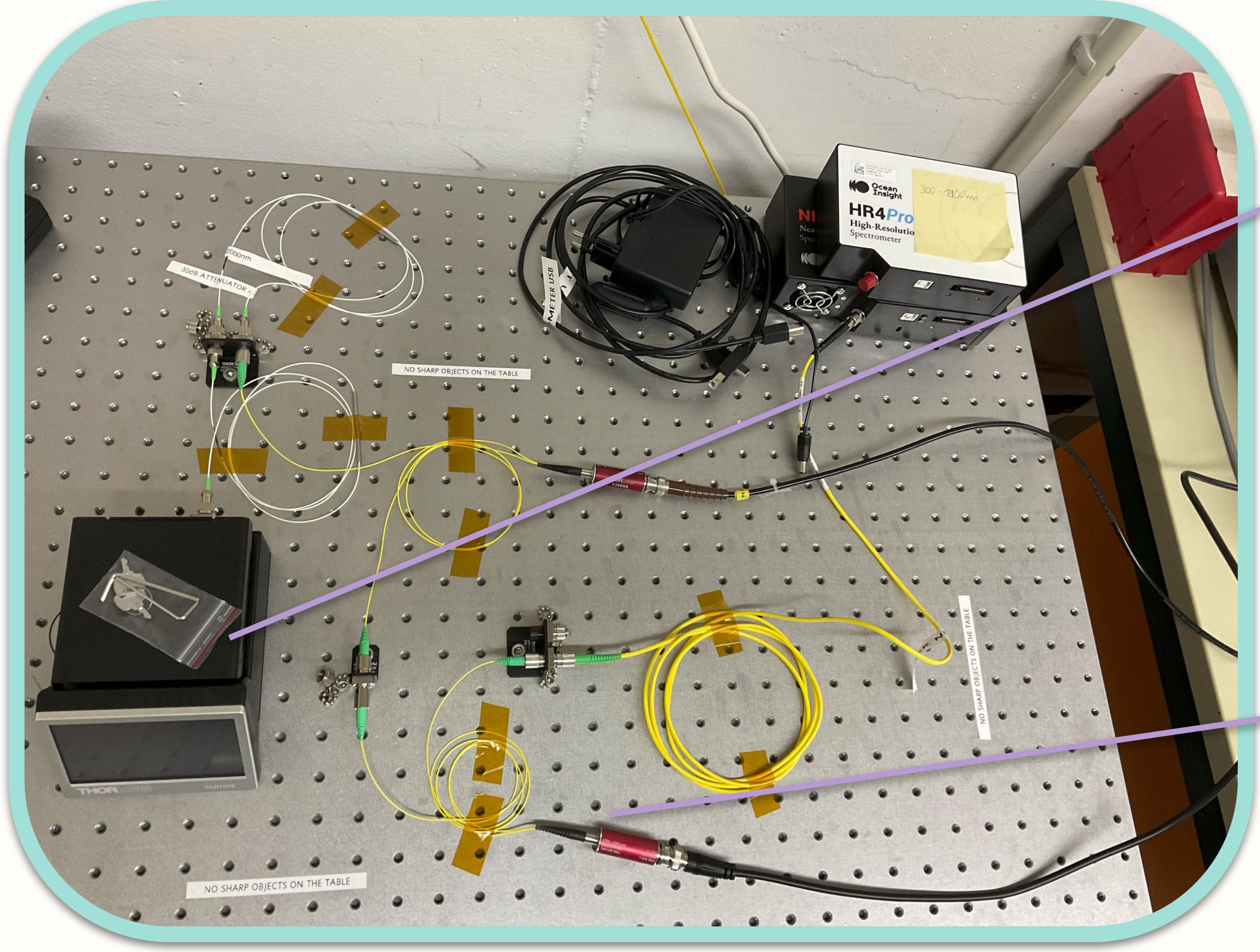


N. Bastidon, The cryogenic photon detection system for the ALPS II experiment: characterization, optimization and background reduction, PhD Thesis, Universität Hamburg, 2016.

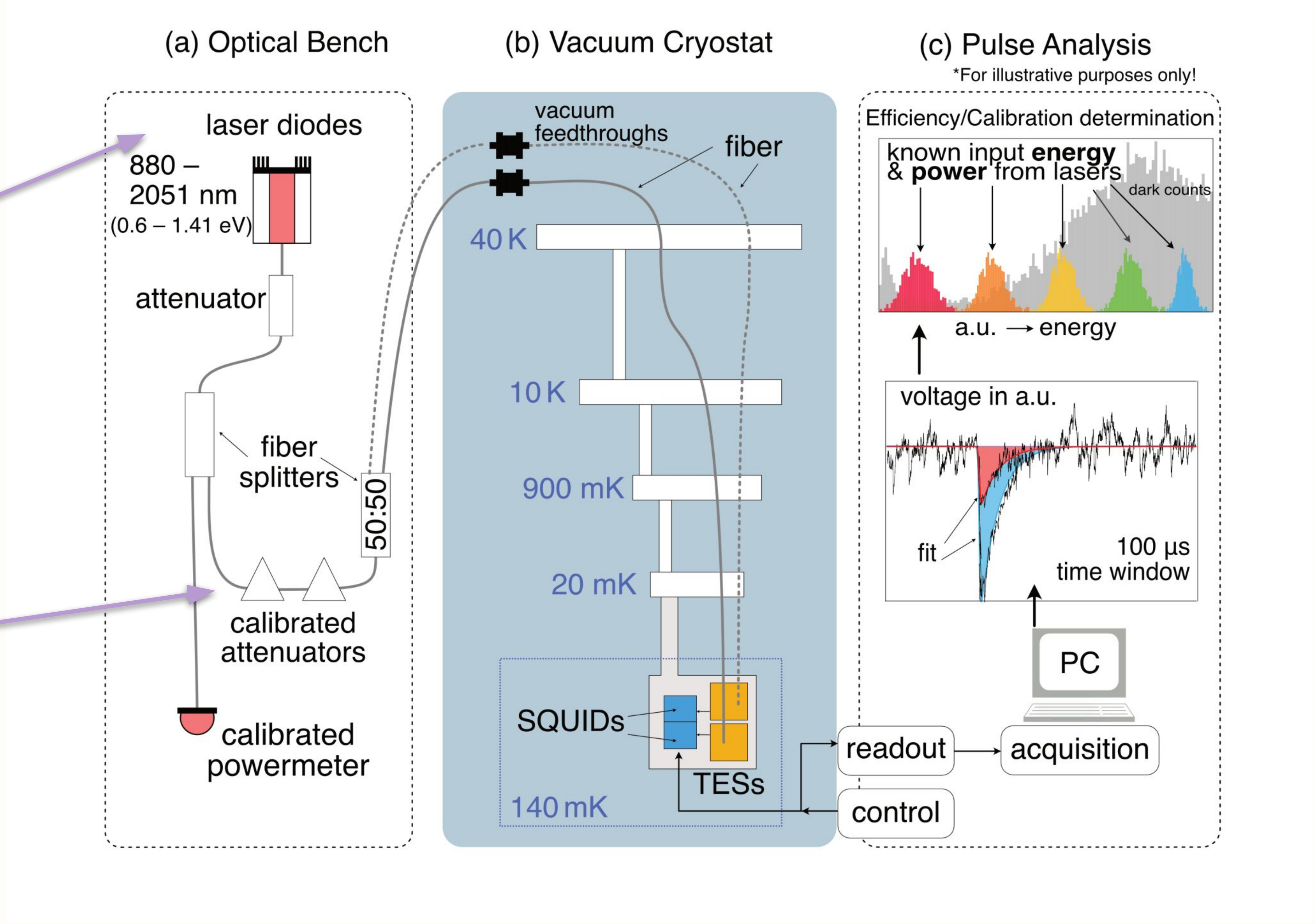


Simulated pulses of different wavelengths with fits (by Jose A. Rubiera Gimeno)

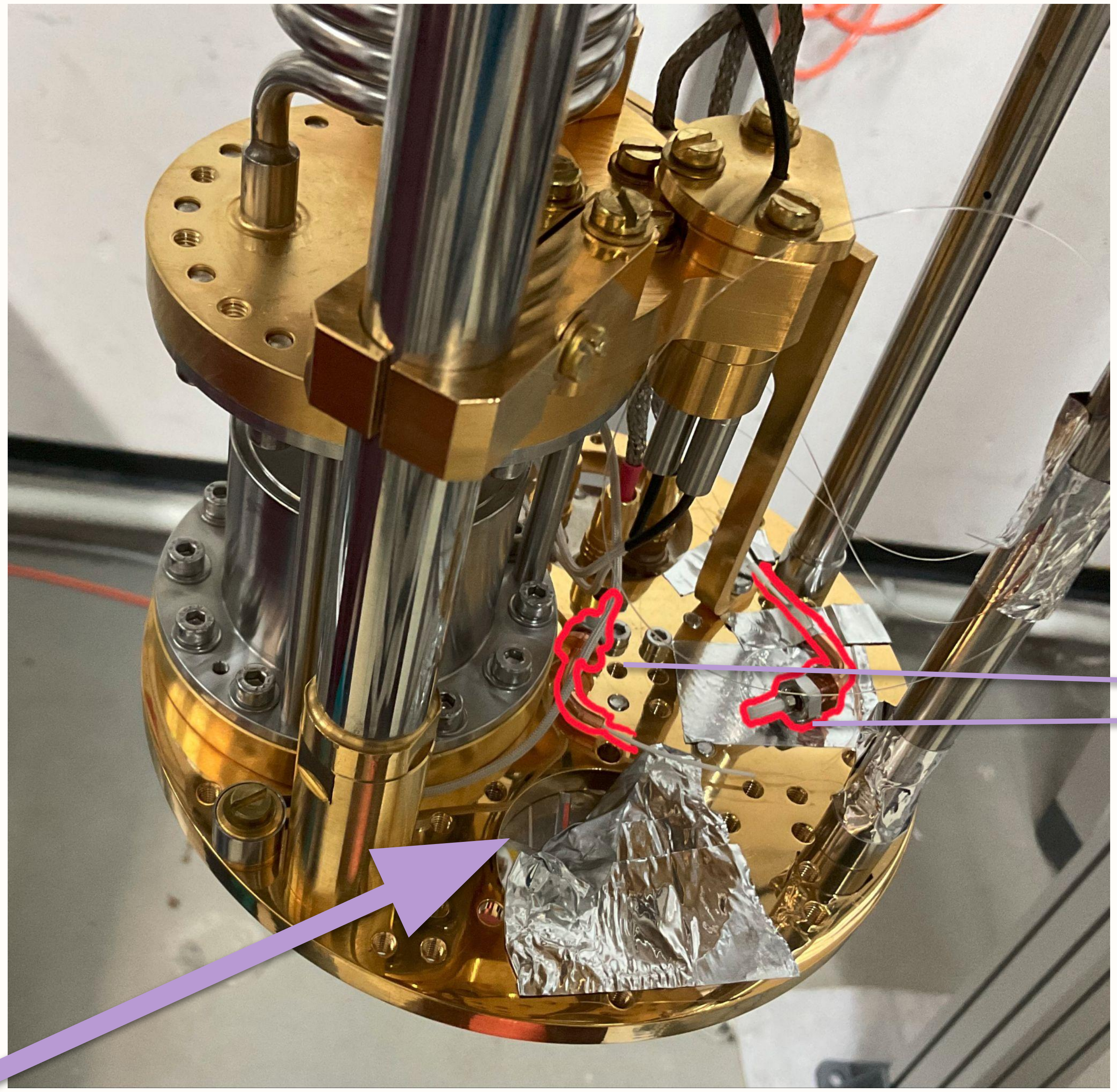
Calibration Setup



Experimental Setup including laser diodes

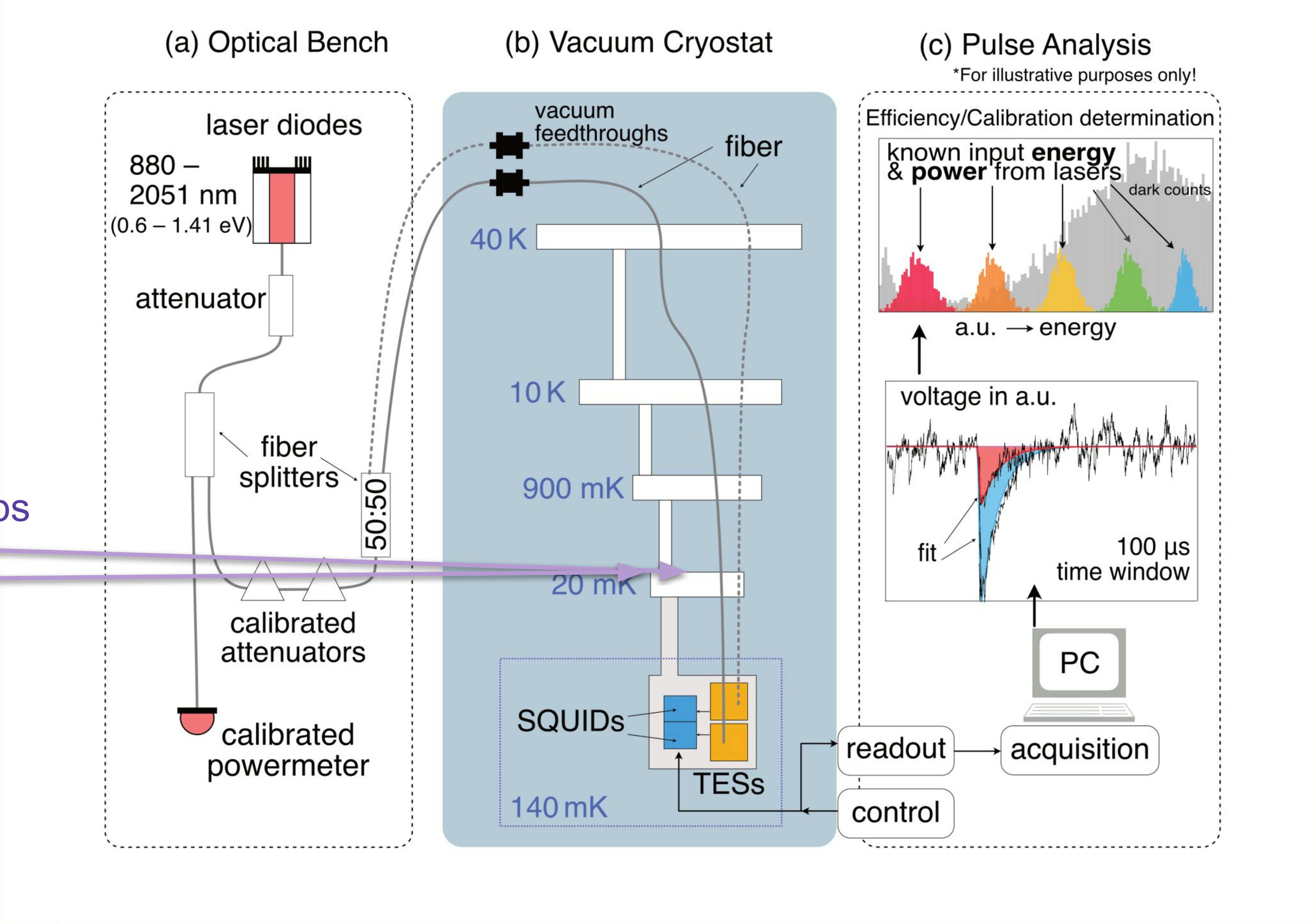


Calibration Setup



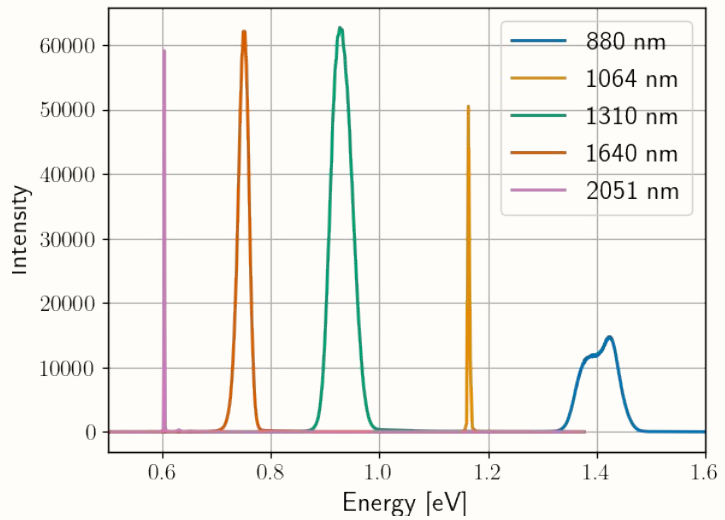
fiber tips

slit

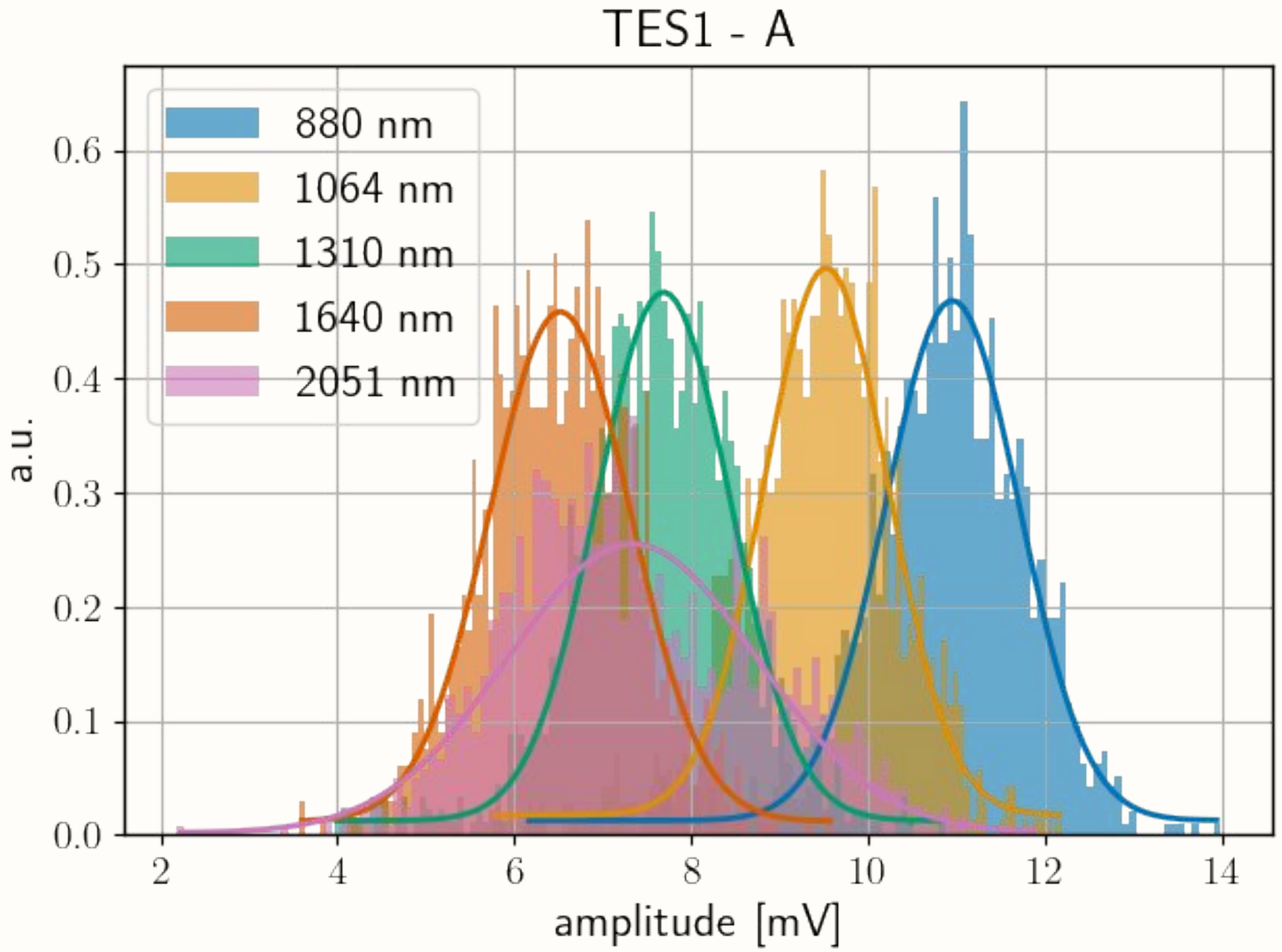


To mimic an intrinsic measurement, the fibers are placed on top of the 20 mK stage above the TES module to reduce black body radiation from the warmer parts of the cryostat. The light can enter the TES space through a small slit.

Calibration Results

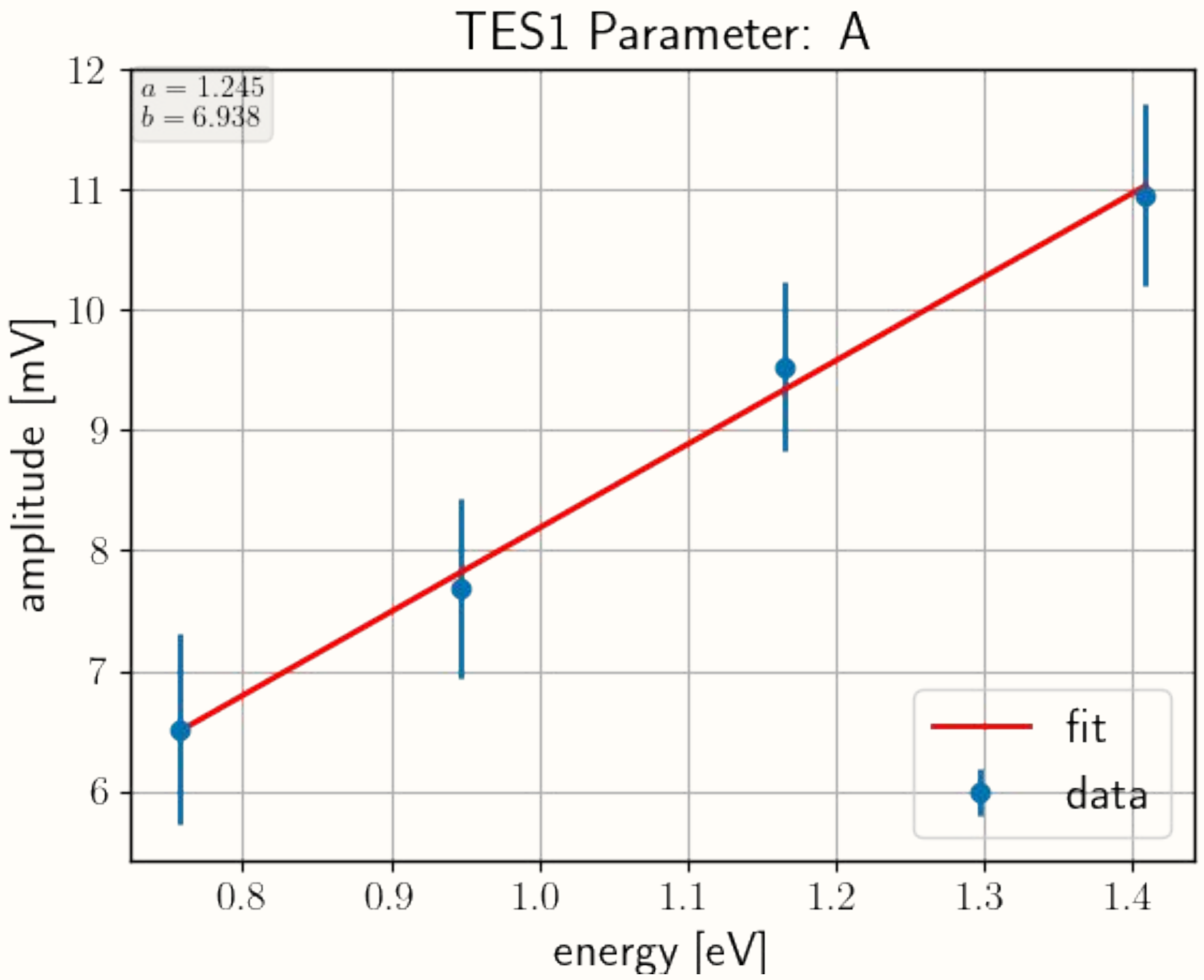


Energy spectrum of used laser diodes



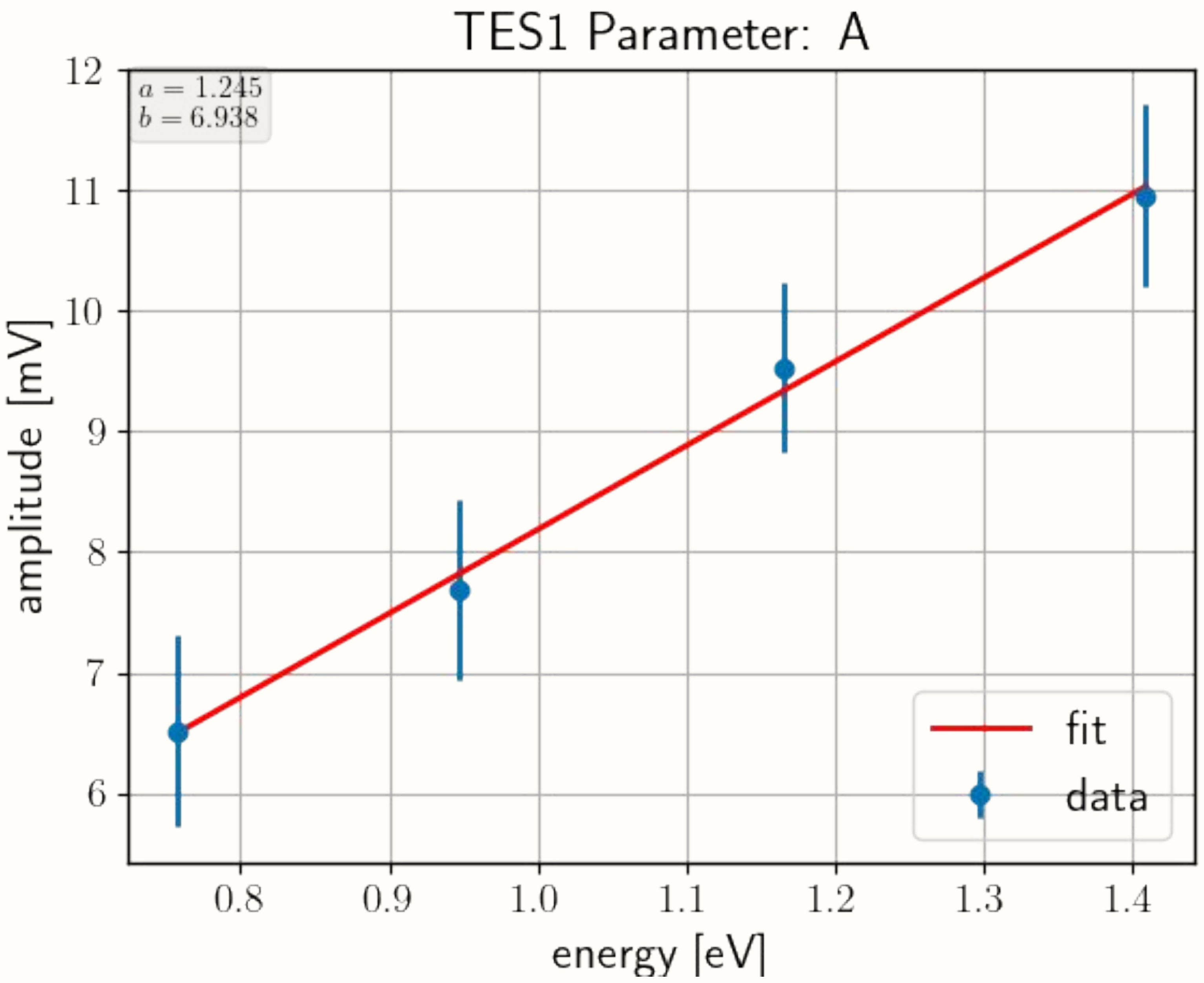
Fit parameters for different wavelengths

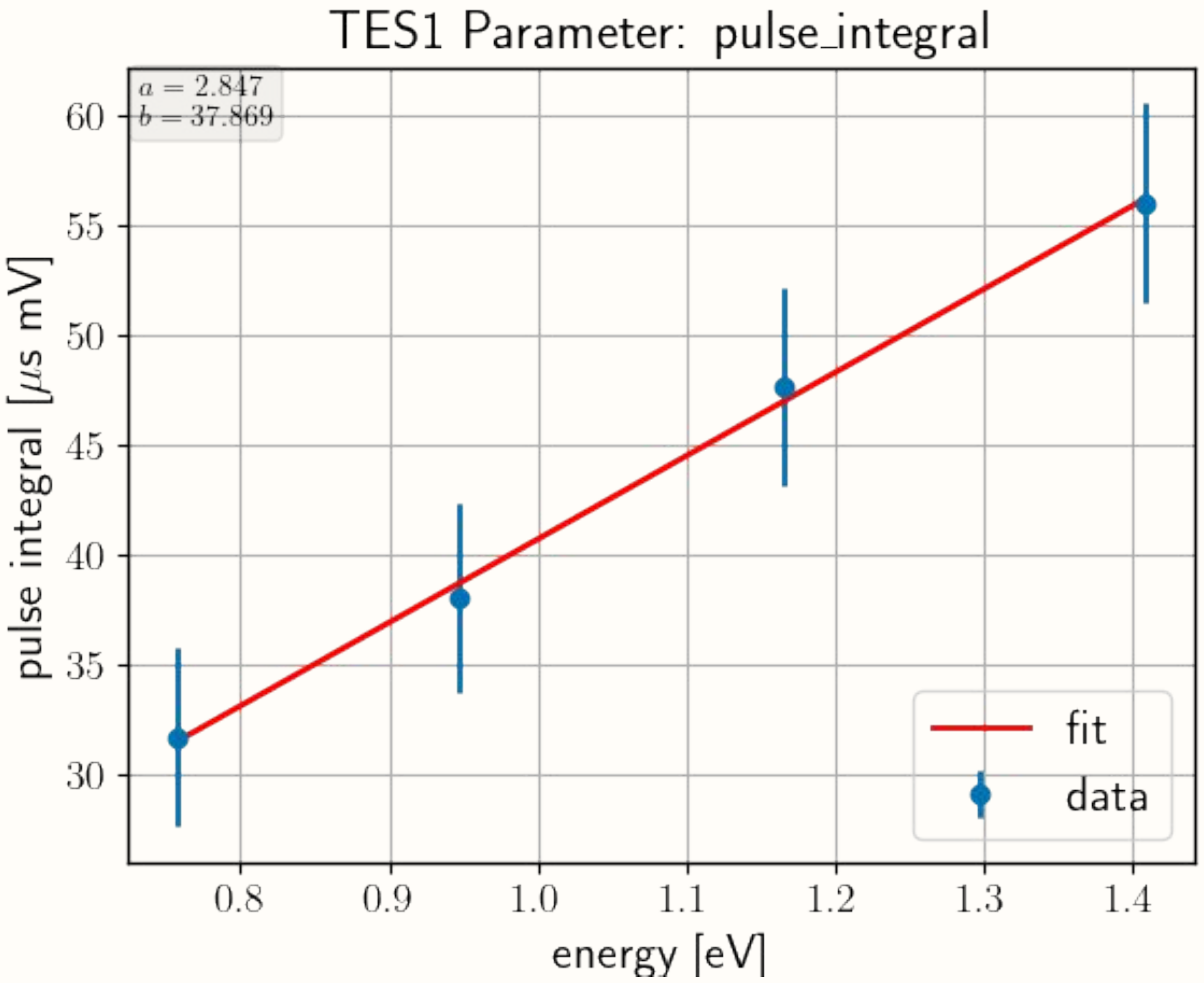
Calibration curve for pulse amplitude



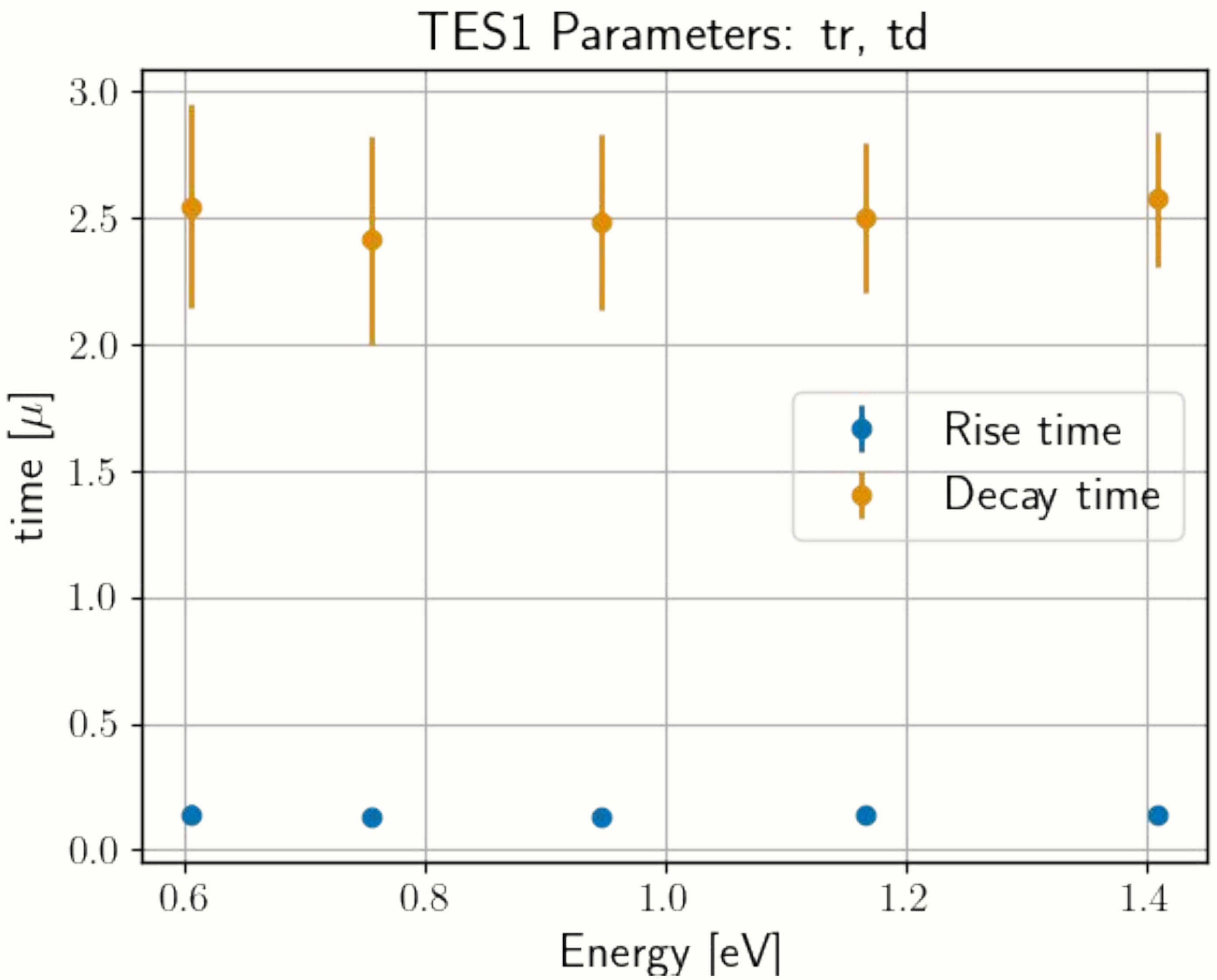
Calibration Results

Amplitude and pulse integral show linearity for the same setup and cooldown



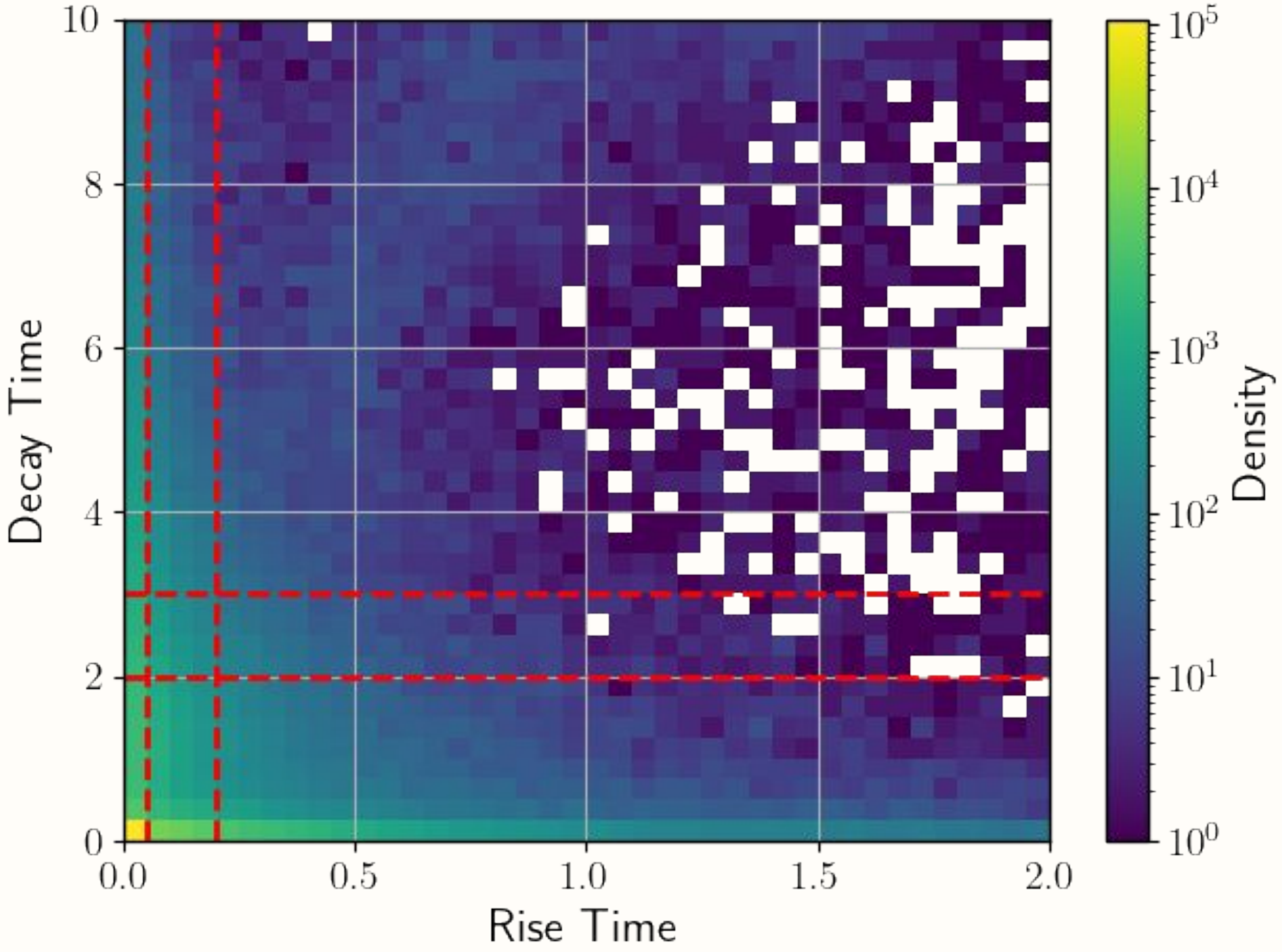


Calibration Results



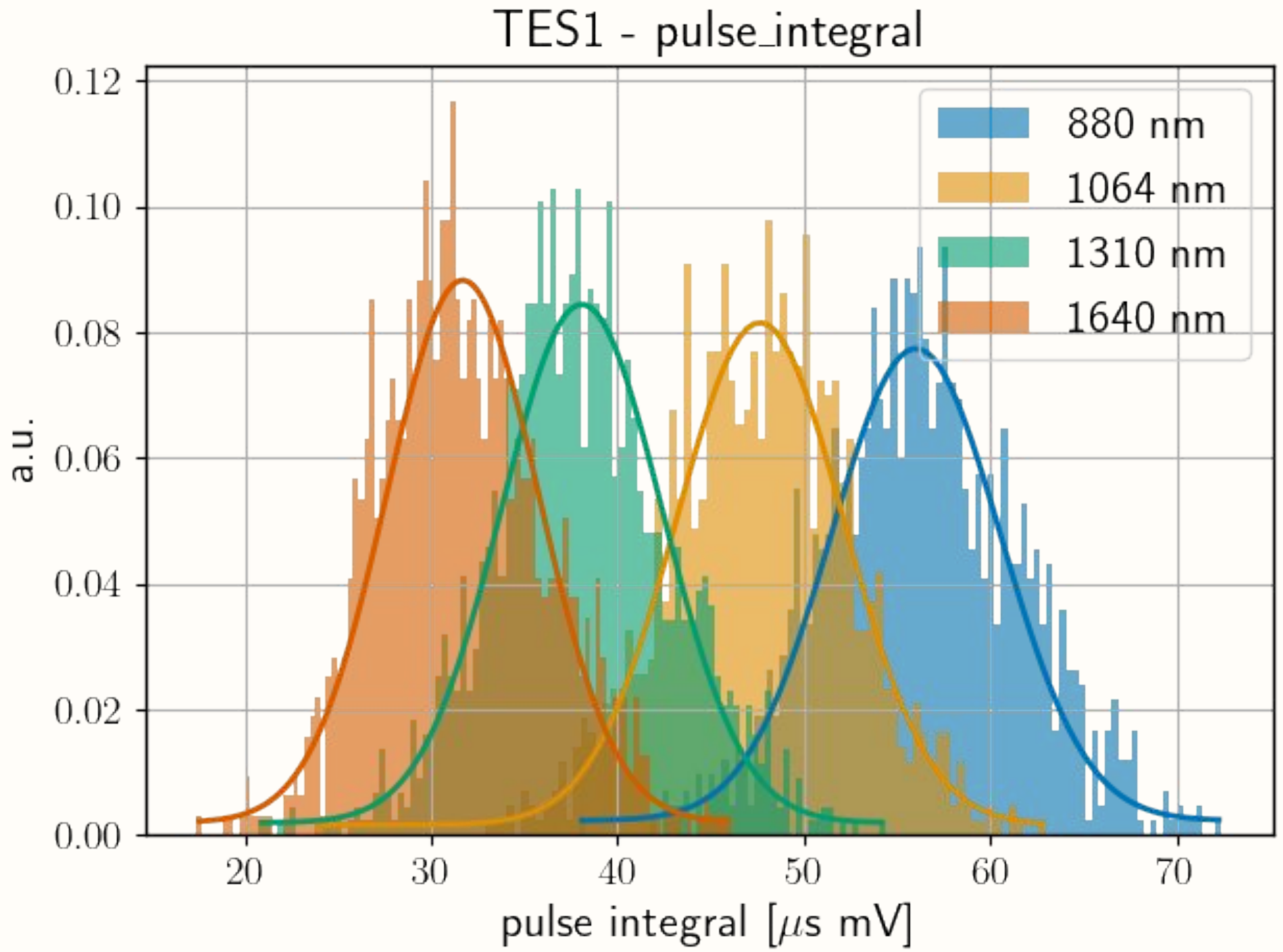
Rise and decay time stay approximately constant over the energy range.

Important to eliminate backgrounds for low trigger thresholds
(Here: 66 h background measurement)

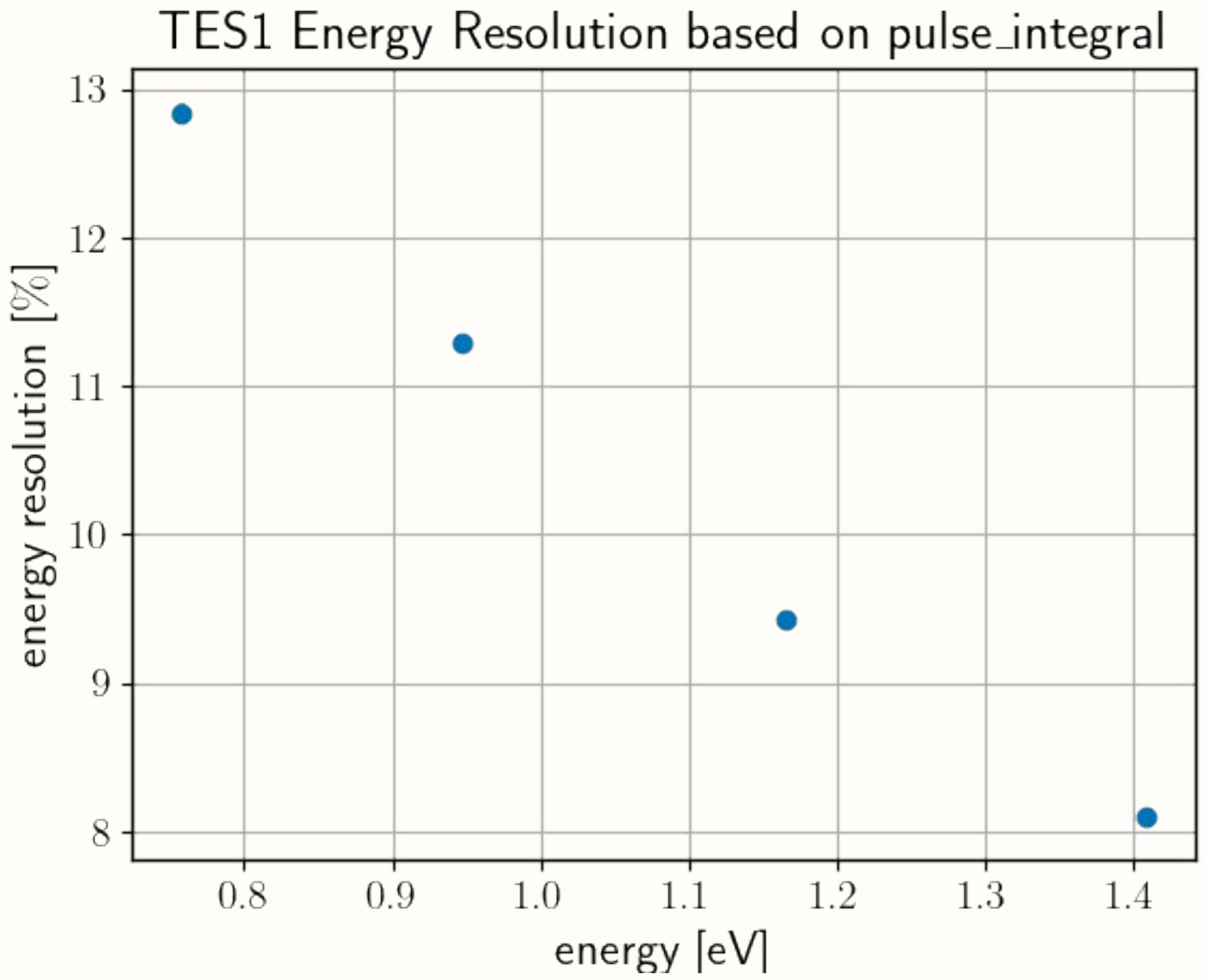


Energy Resolution

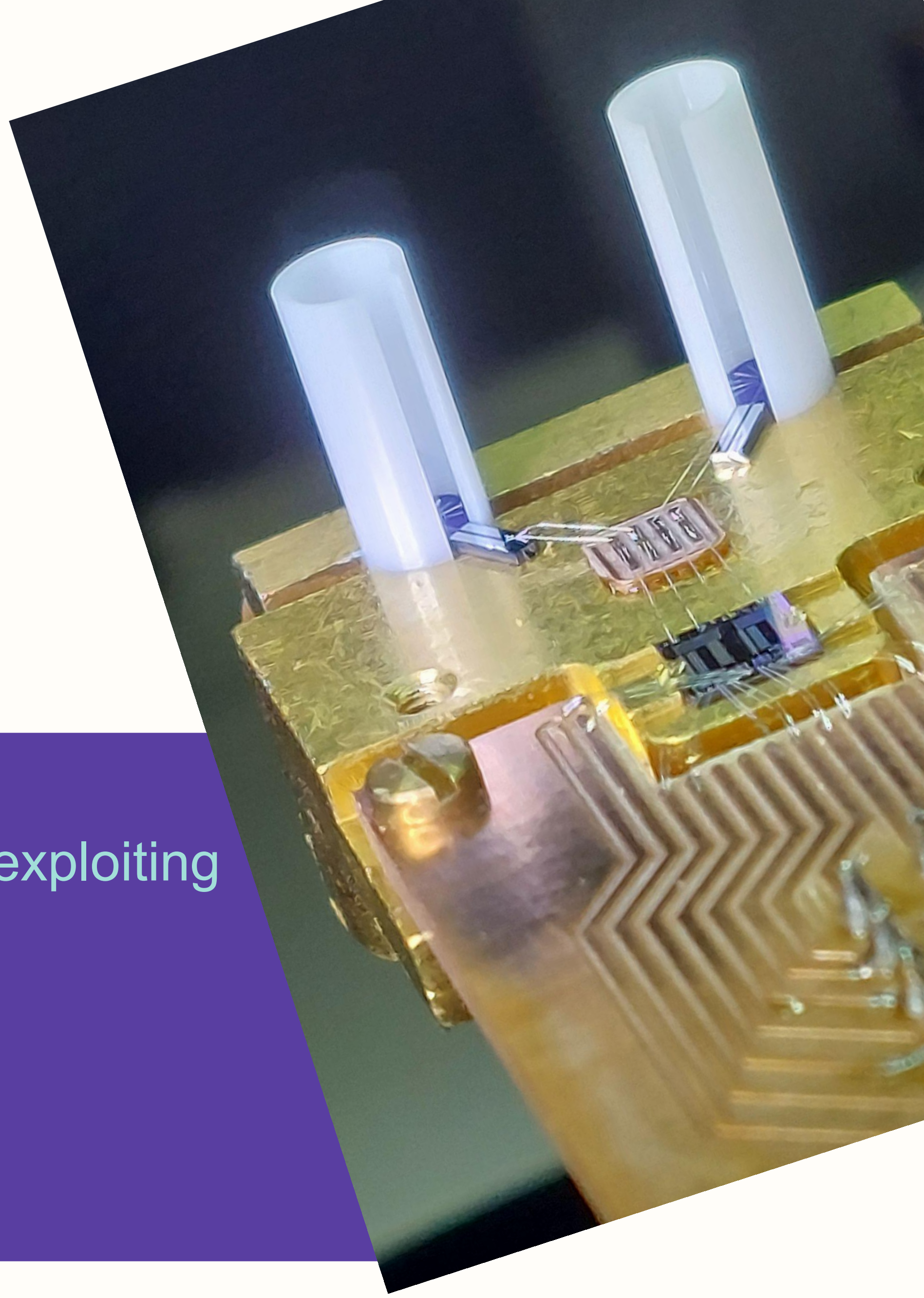
Energy resolution improves for higher energies
(expected due to higher noise contribution)



$$ER = \frac{\sigma}{\mu}$$

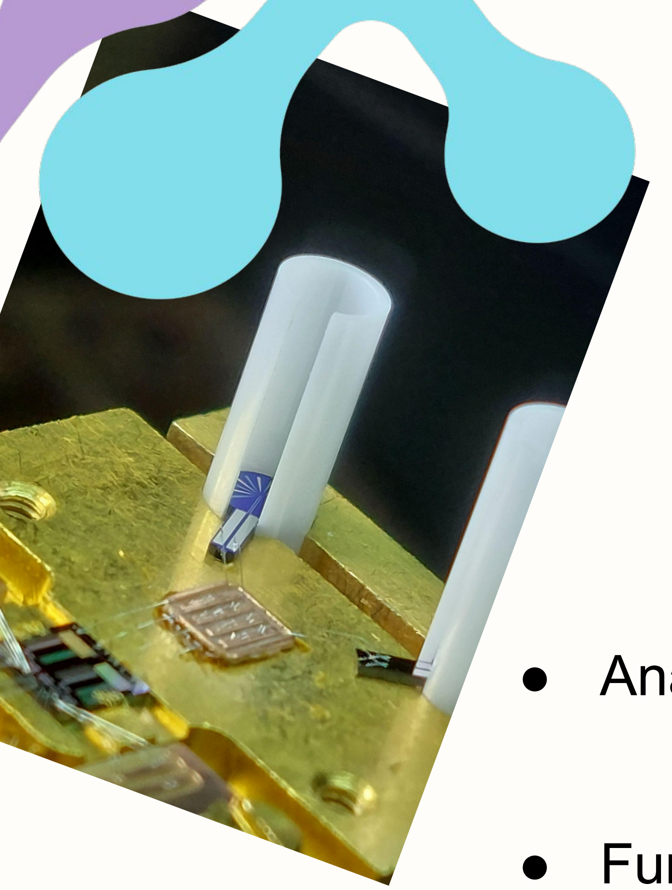


Summary

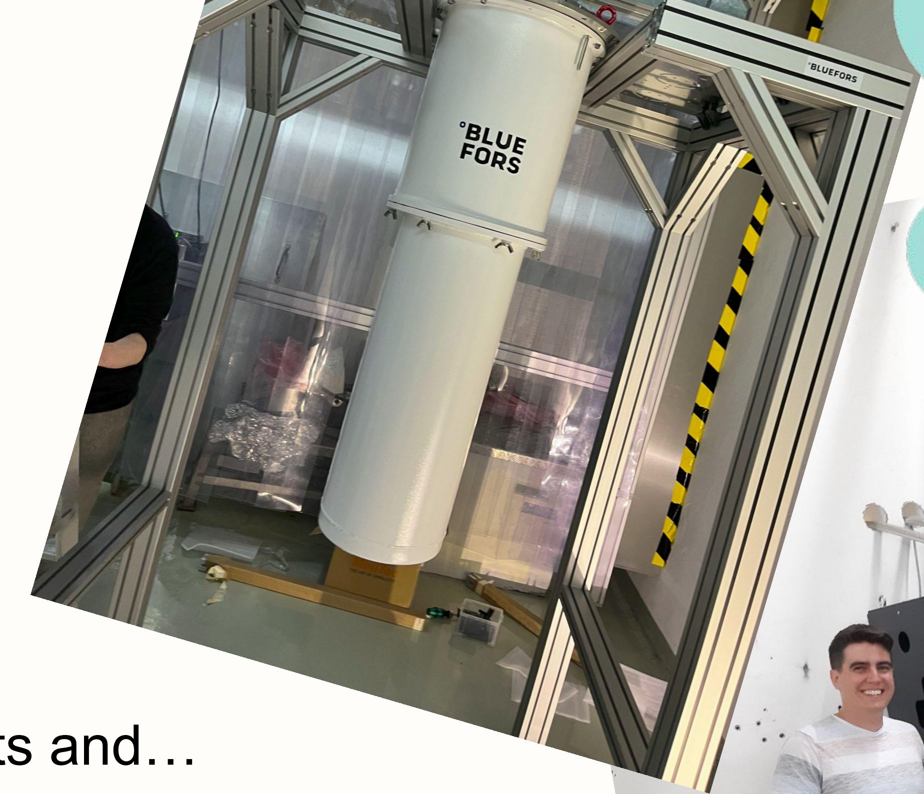


- Our TES can search for sub-MeV DM prior to being employed in the ALPS II experiment exploiting DM-electron scattering possibly reaching new sensitivities
- This could be a proof of principle for similar technologies being used as DM detectors
- 20 day dedicated DDM run performed in April/May, analysis ongoing

- calibration results suggest linear relationship between pulse amplitude/integral and energy of the incoming particle
- rise and decay time stay constant over a large range of energies
  → calibration enables a better understanding and supports the analysis of dedicated direct DM searches



Outlook



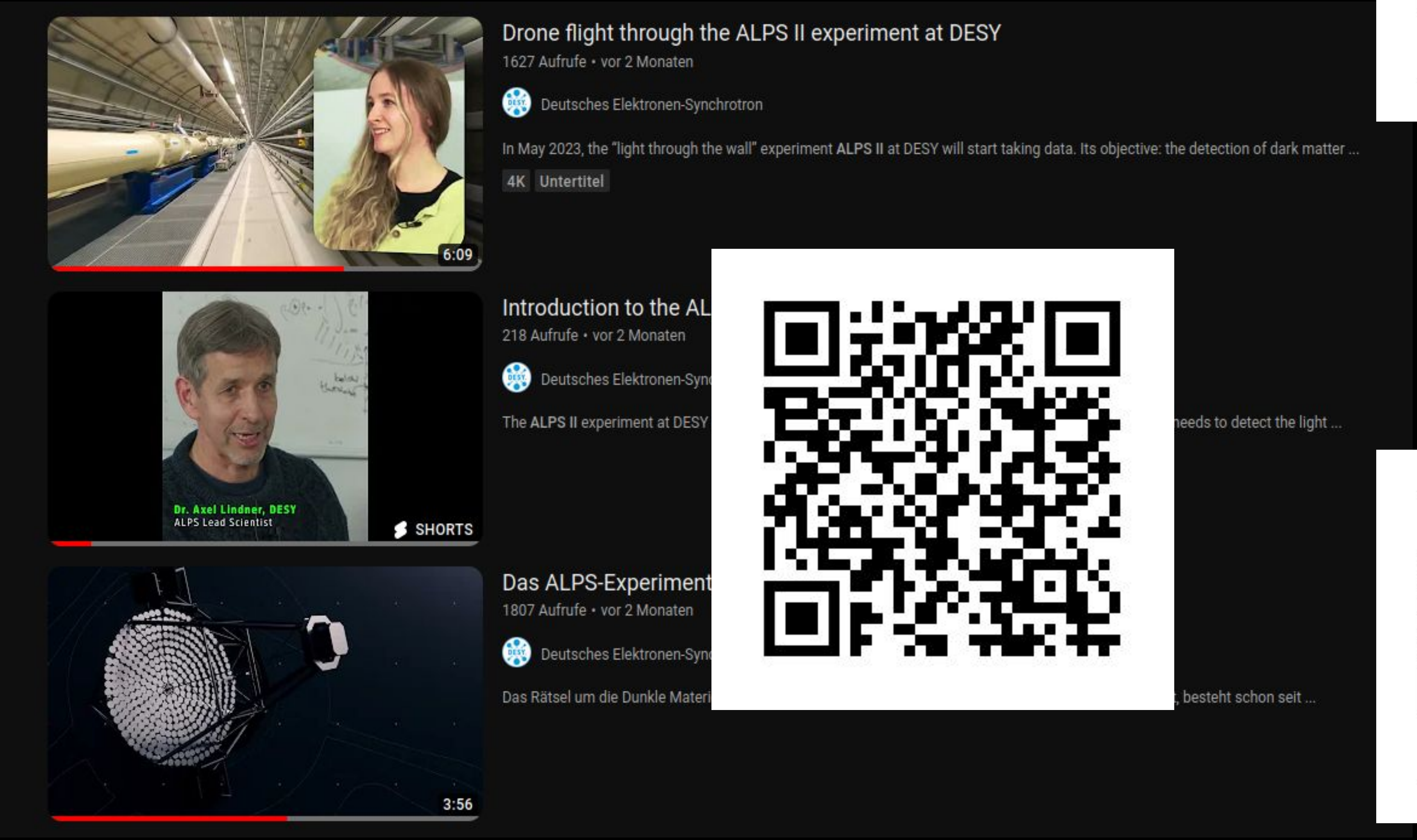
- Analysis of dedicated DM run ongoing - supported by calibration results and…
- Further investigation of intrinsic background (paper soon: *Simulating the response of a transition edge sensor to cosmic-ray and radioactivity induced backgrounds* by J.A. Rubiera Gimeno et al.)
- Expecting to receive 1550 nm TES from NIST without zirconia fiber sleeves that could reduce intrinsic backgrounds (see paper!)
- further investigate behavior < 0.8 eV and repeat measurements with additional sensors (e.g. with different SQUID chips)

THANK YOU

Animation by SciComLab

More ALPS II

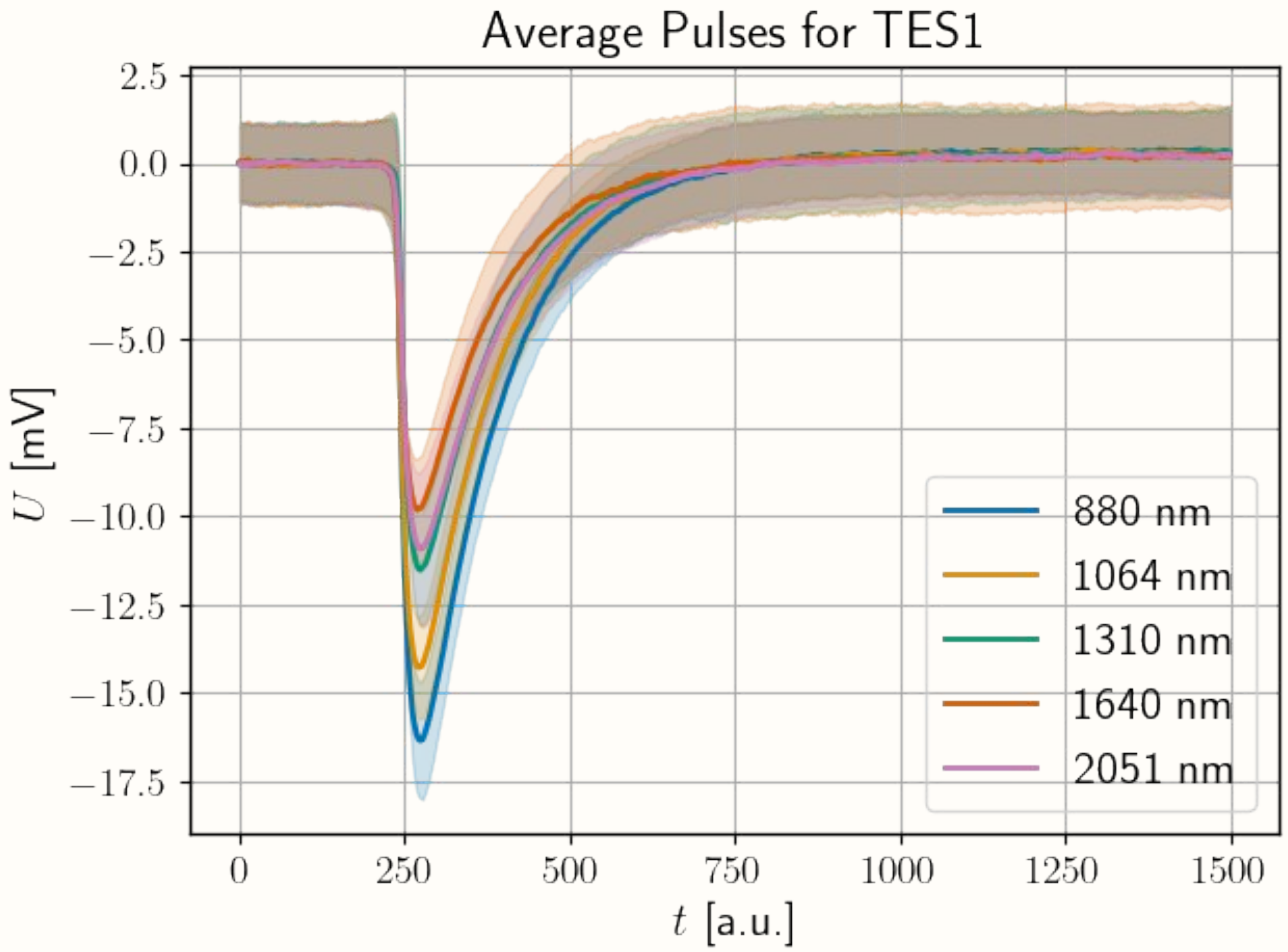
Video material on YOUTUBE:



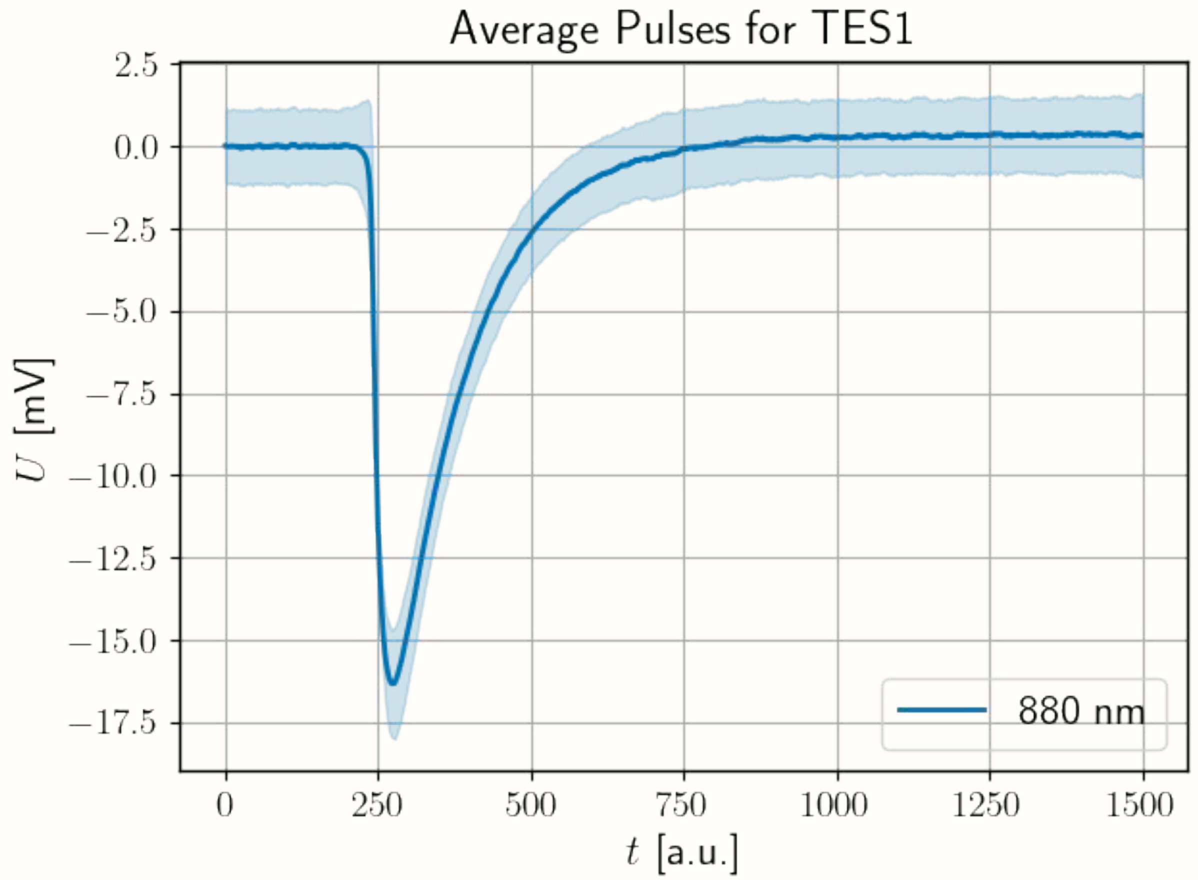
Podcasts (Spotify Links, but also available on other platforms)

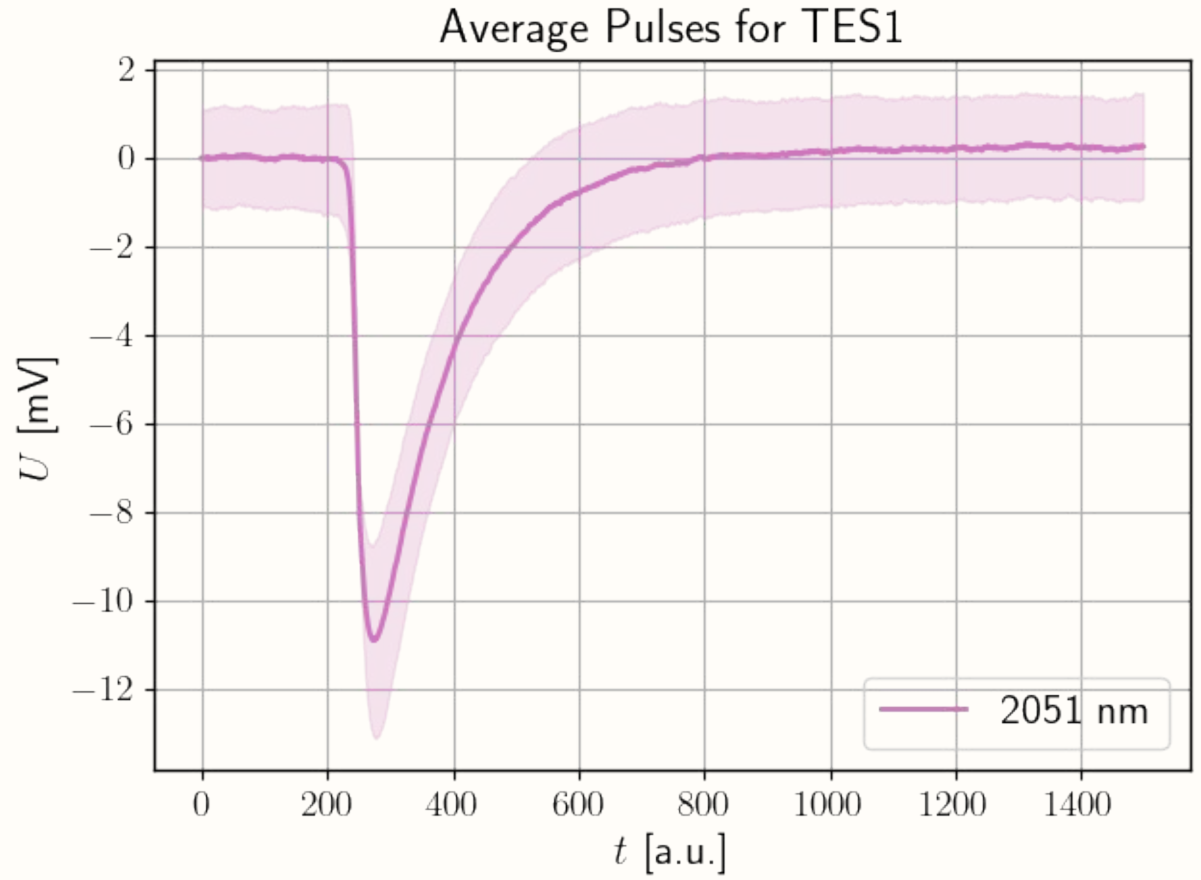
ALPS II @ DESY

Measured pulses



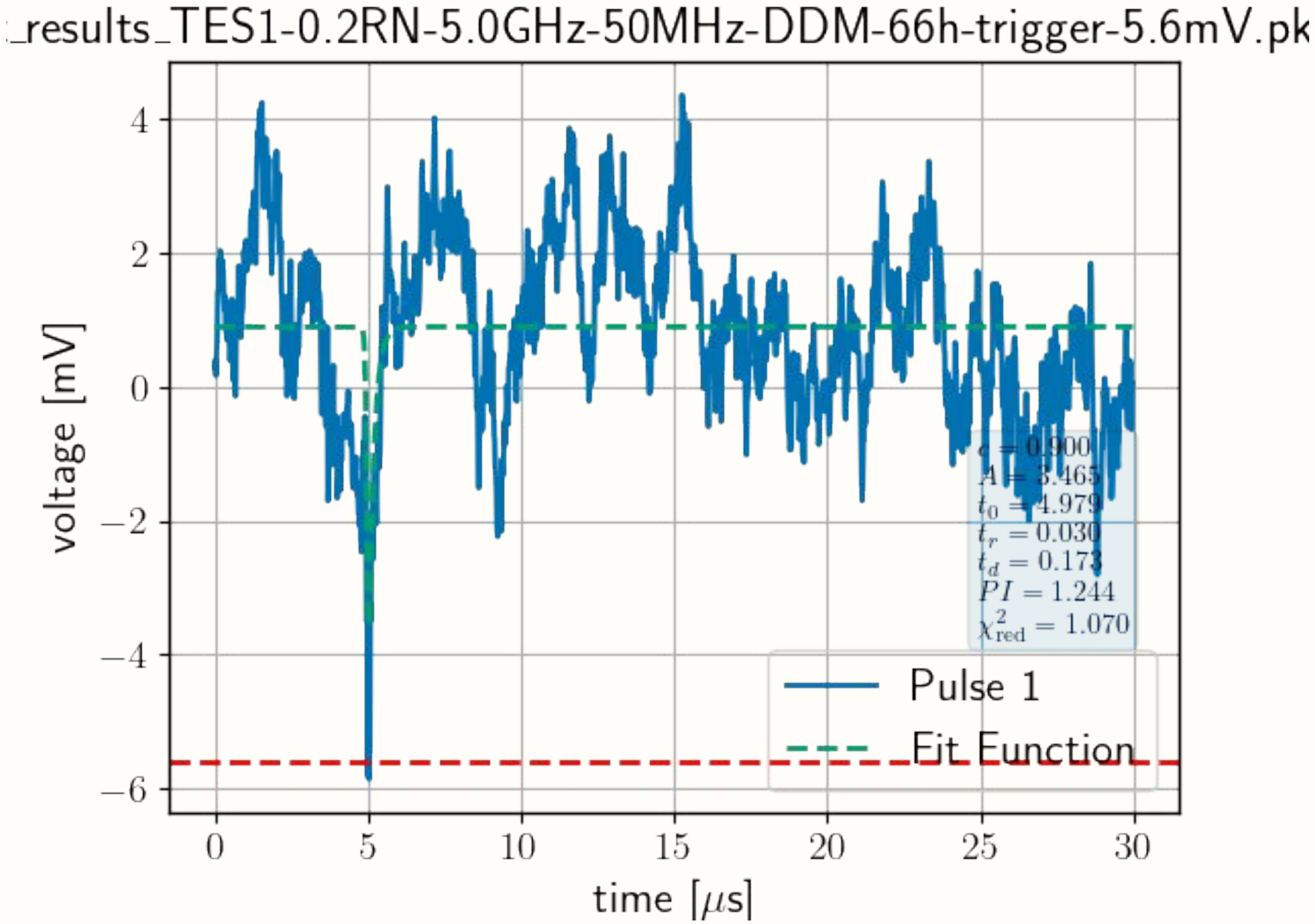
Average pulses of different wavelengths
(calibration measurement)

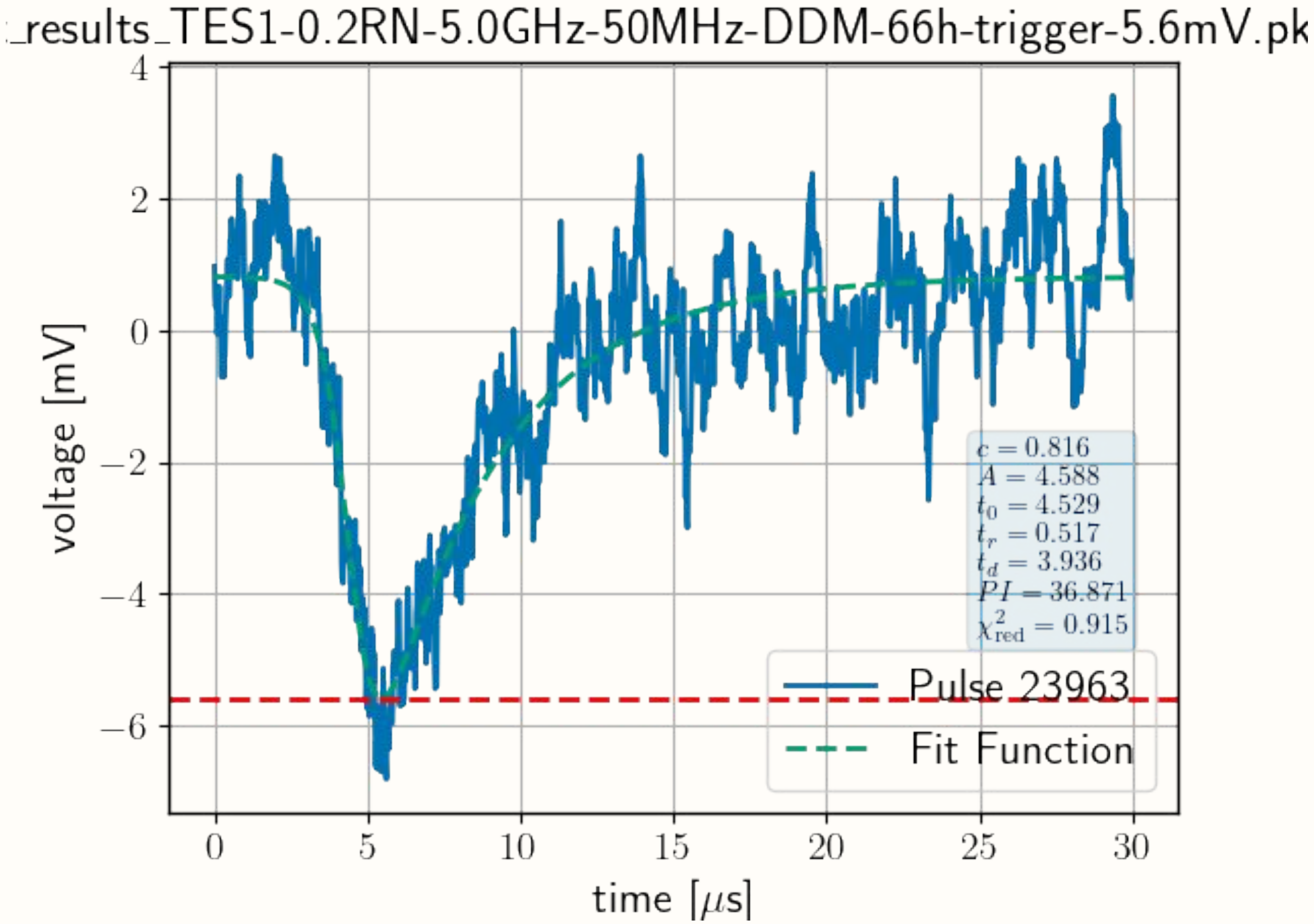




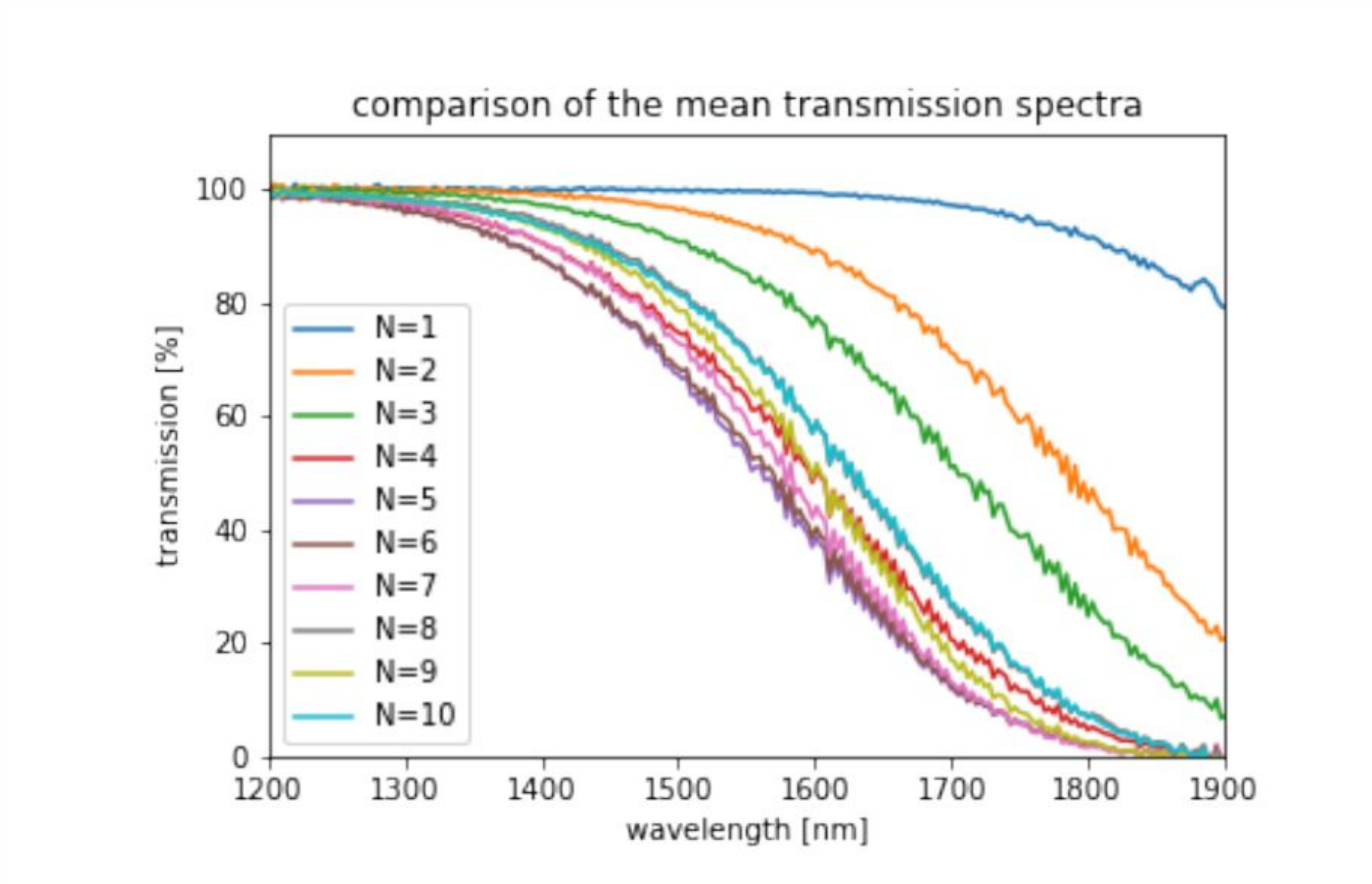
More noisy at lower trigger thresholds

DDM backgrounds - preliminary



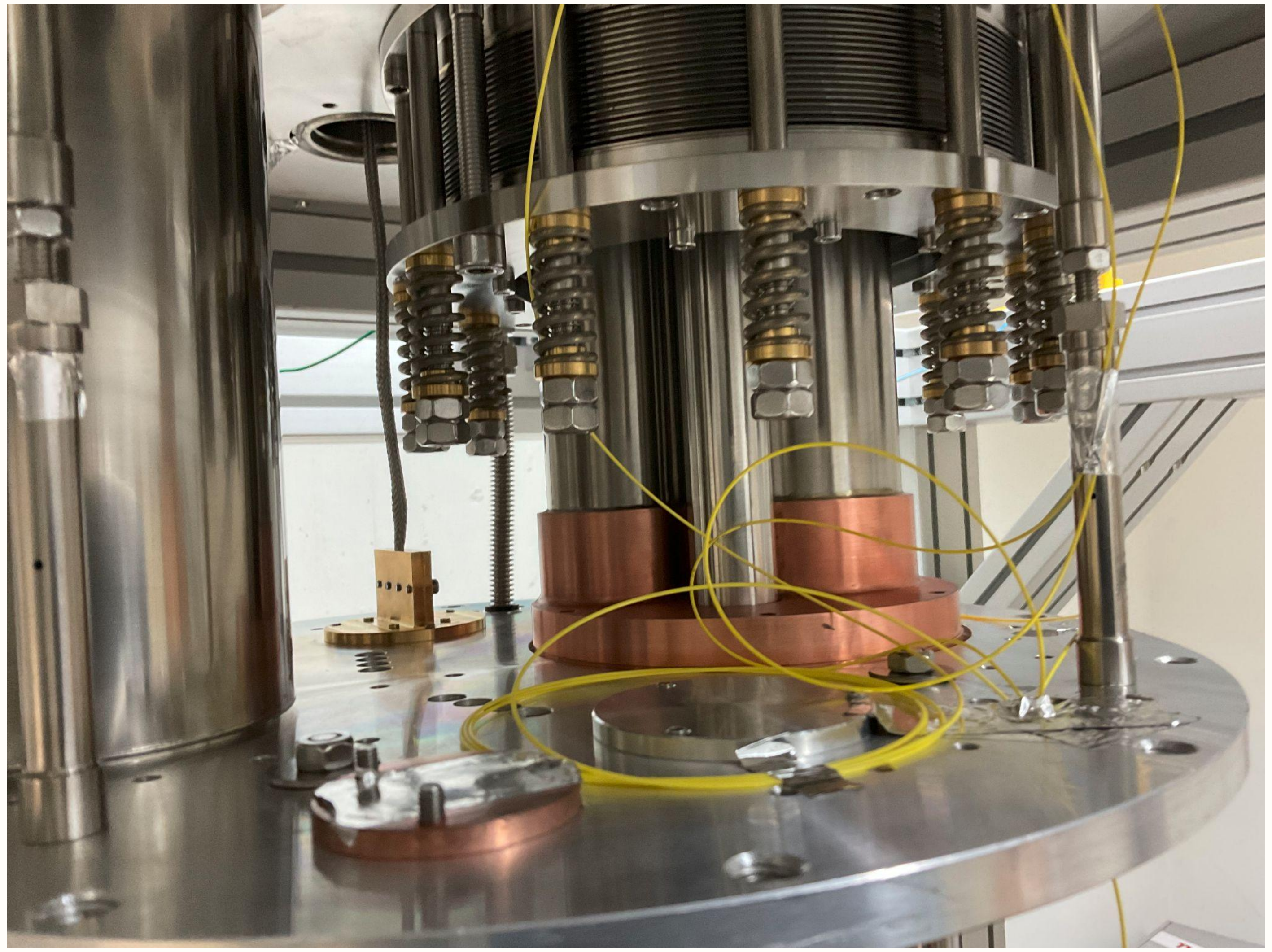


Effects due to fiber curling

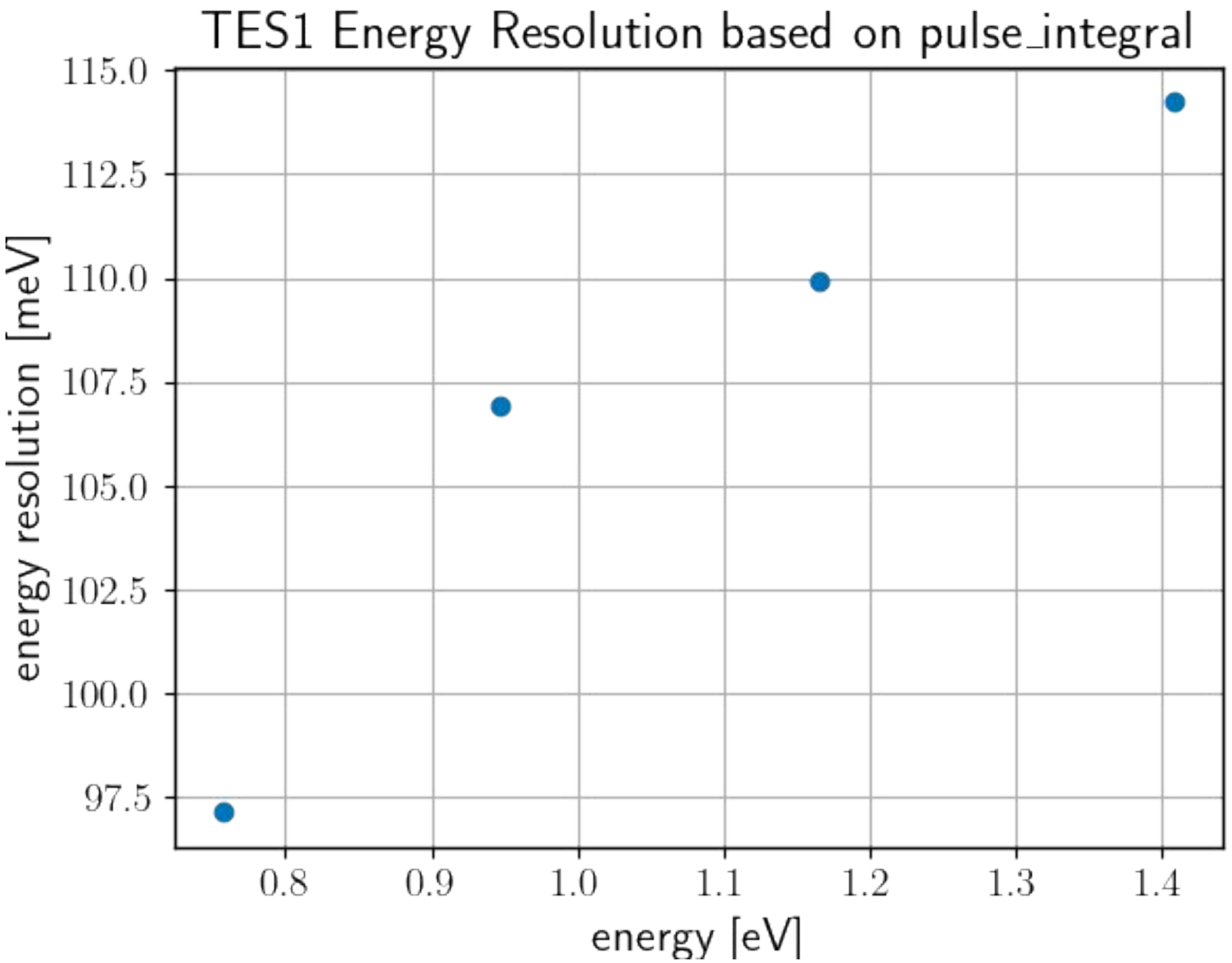


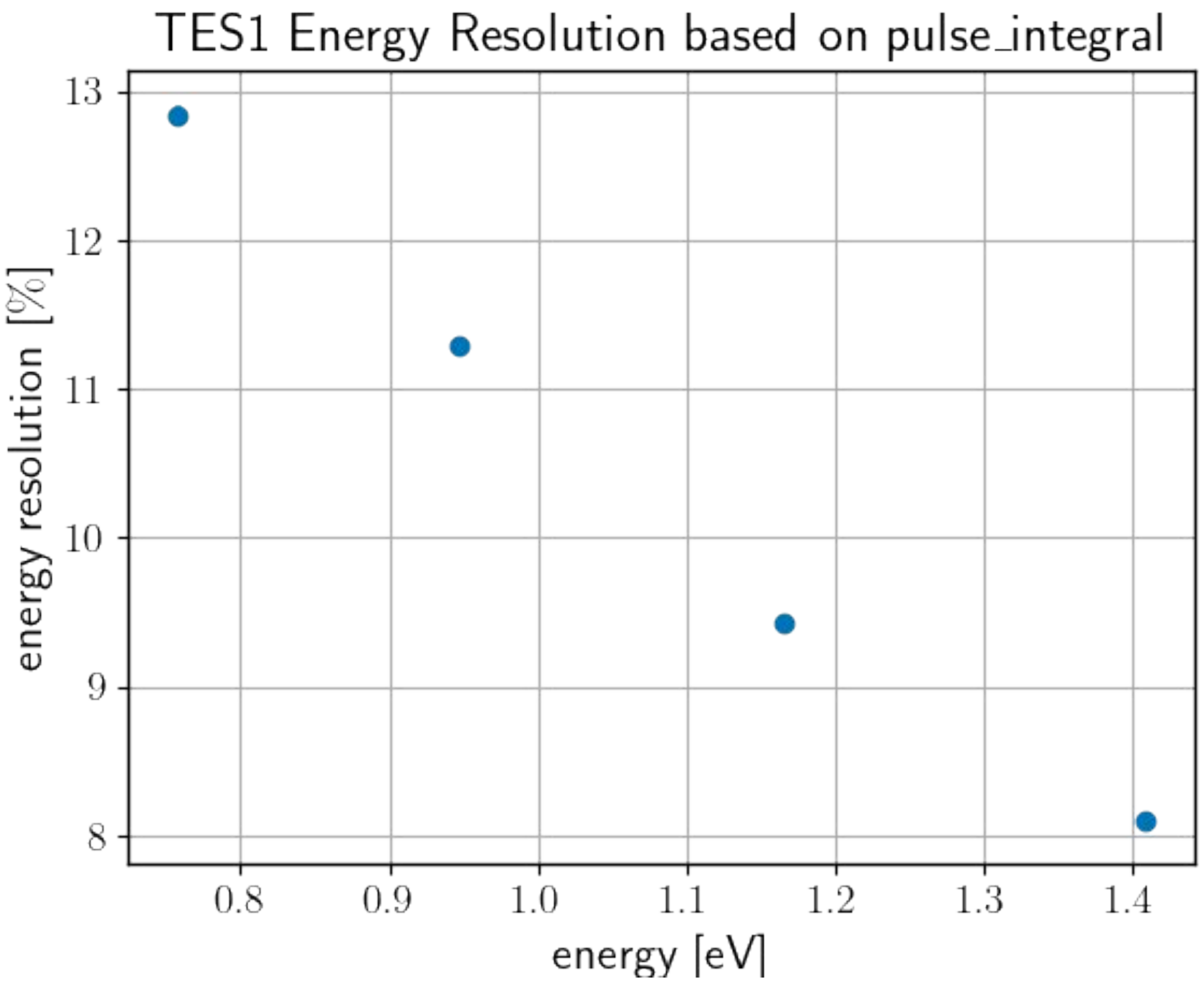
Transmission spectrum of a white light source for different windings of curled fiber - measured by intern Maria T. Pabst

Curled optical fiber on upper part of the cryostat.

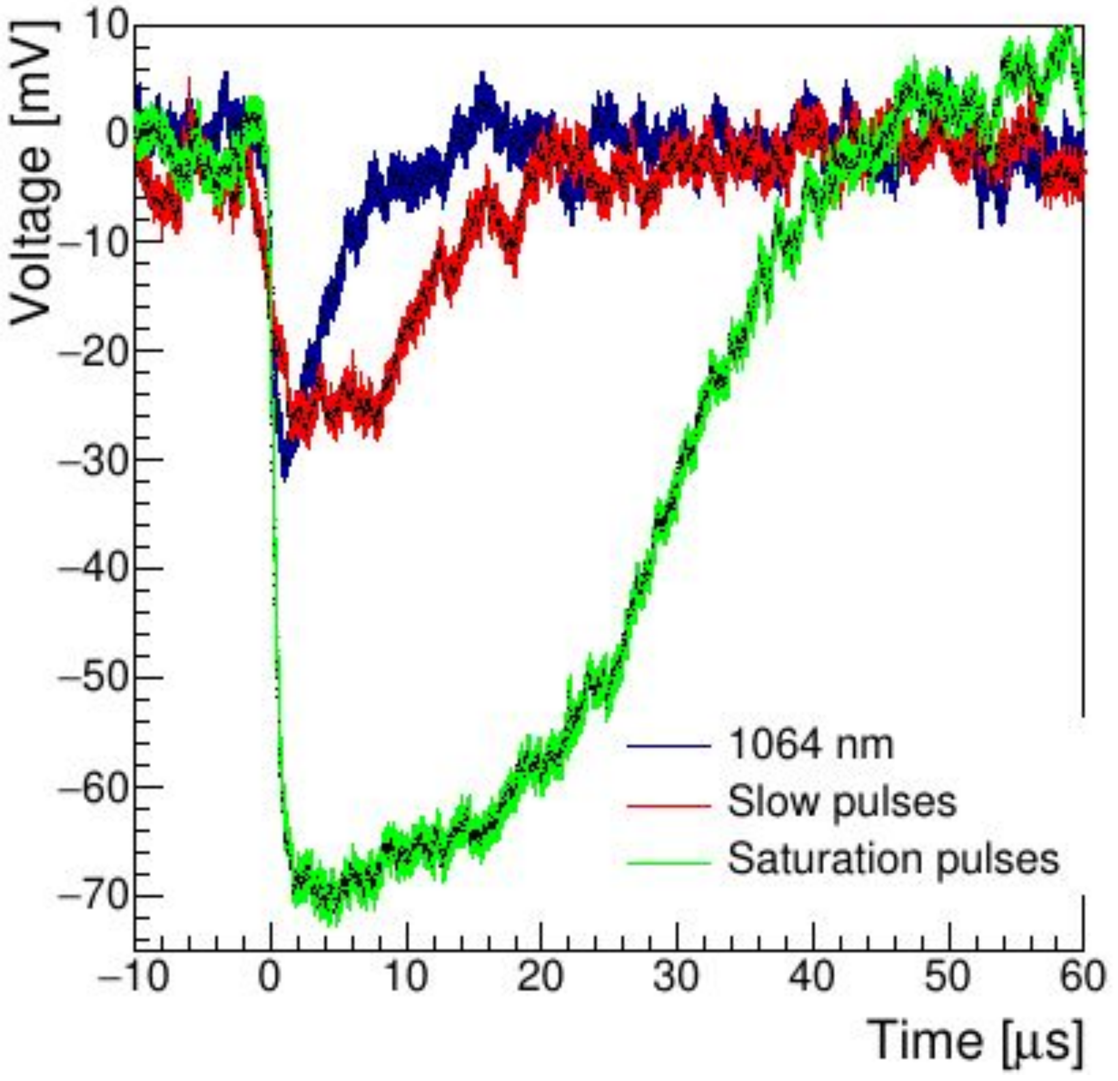


Energy resolution

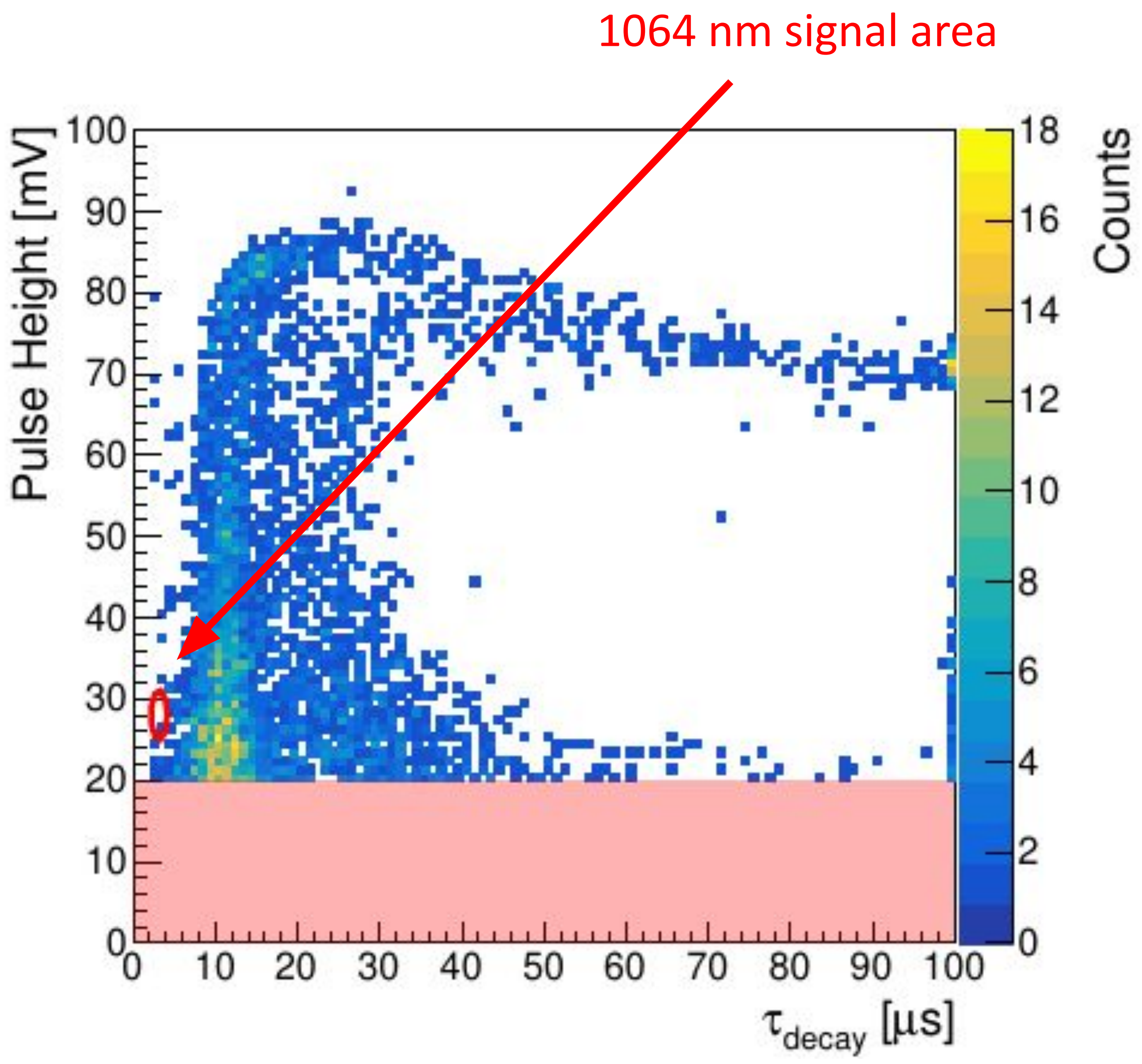




Intrinsic TES Backgrounds



Example pulses for intrinsic background

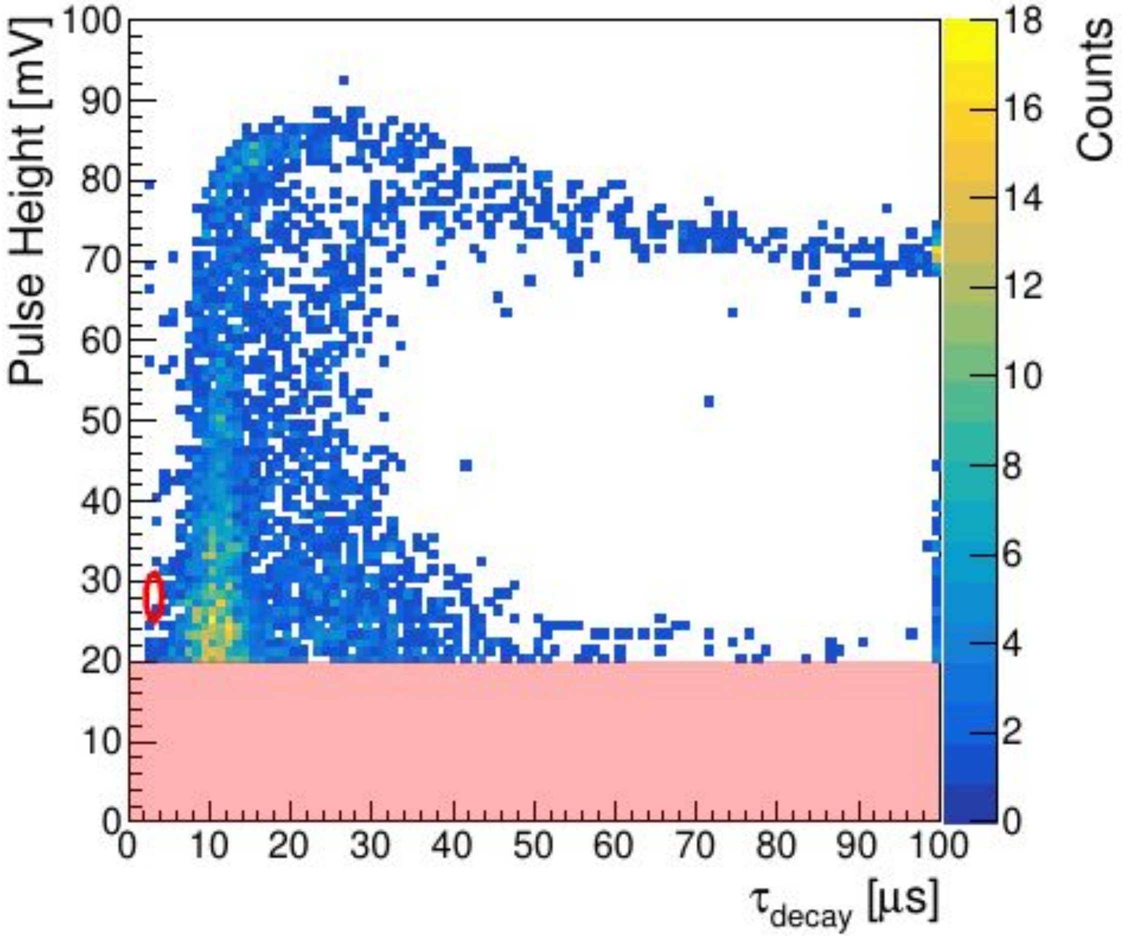


pulse-height distribution of intrinsic background pulses

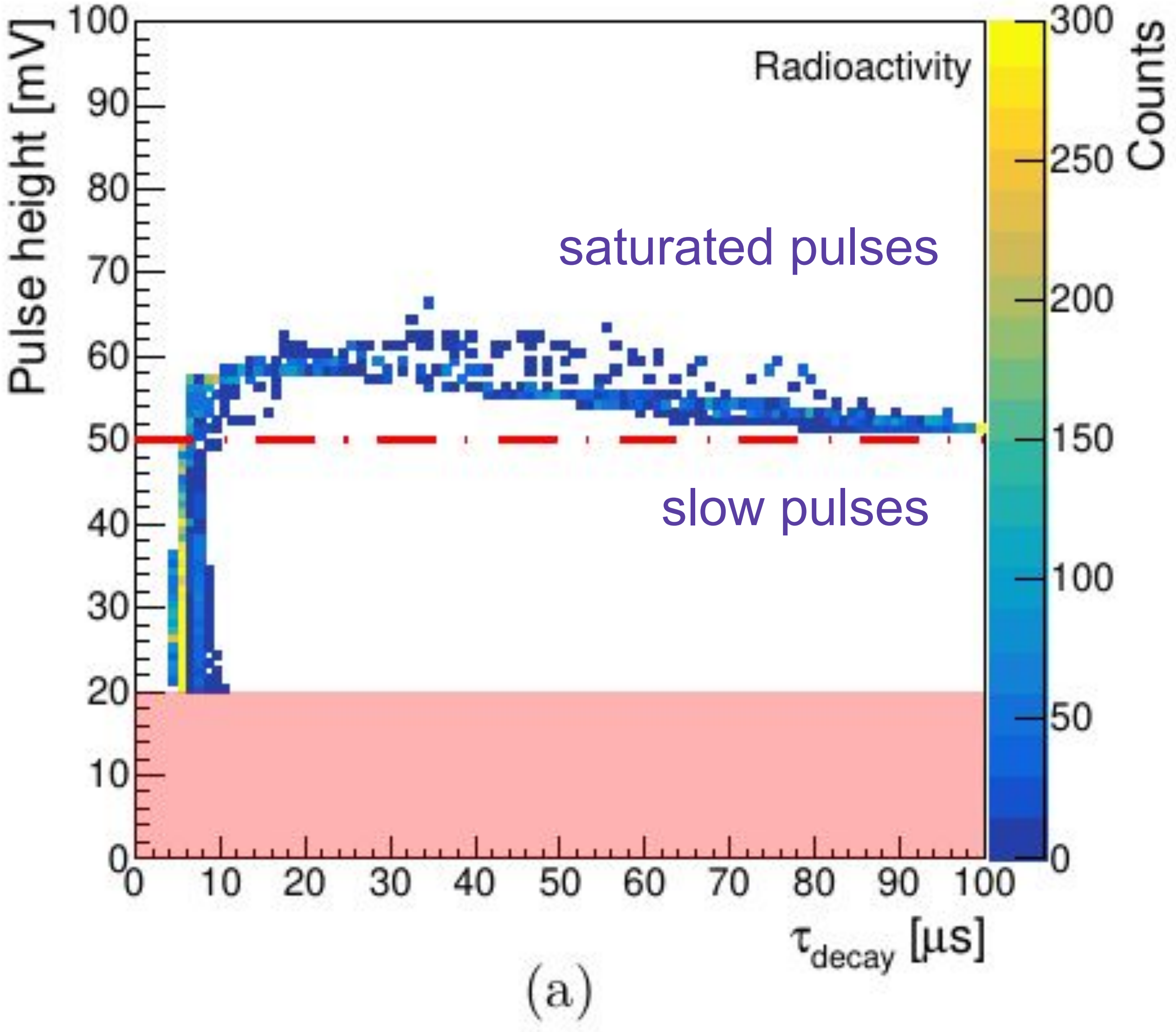
Studies by Jose A. Ruberia Gimeno

Studies by Jose A. Ruberia Gimeno

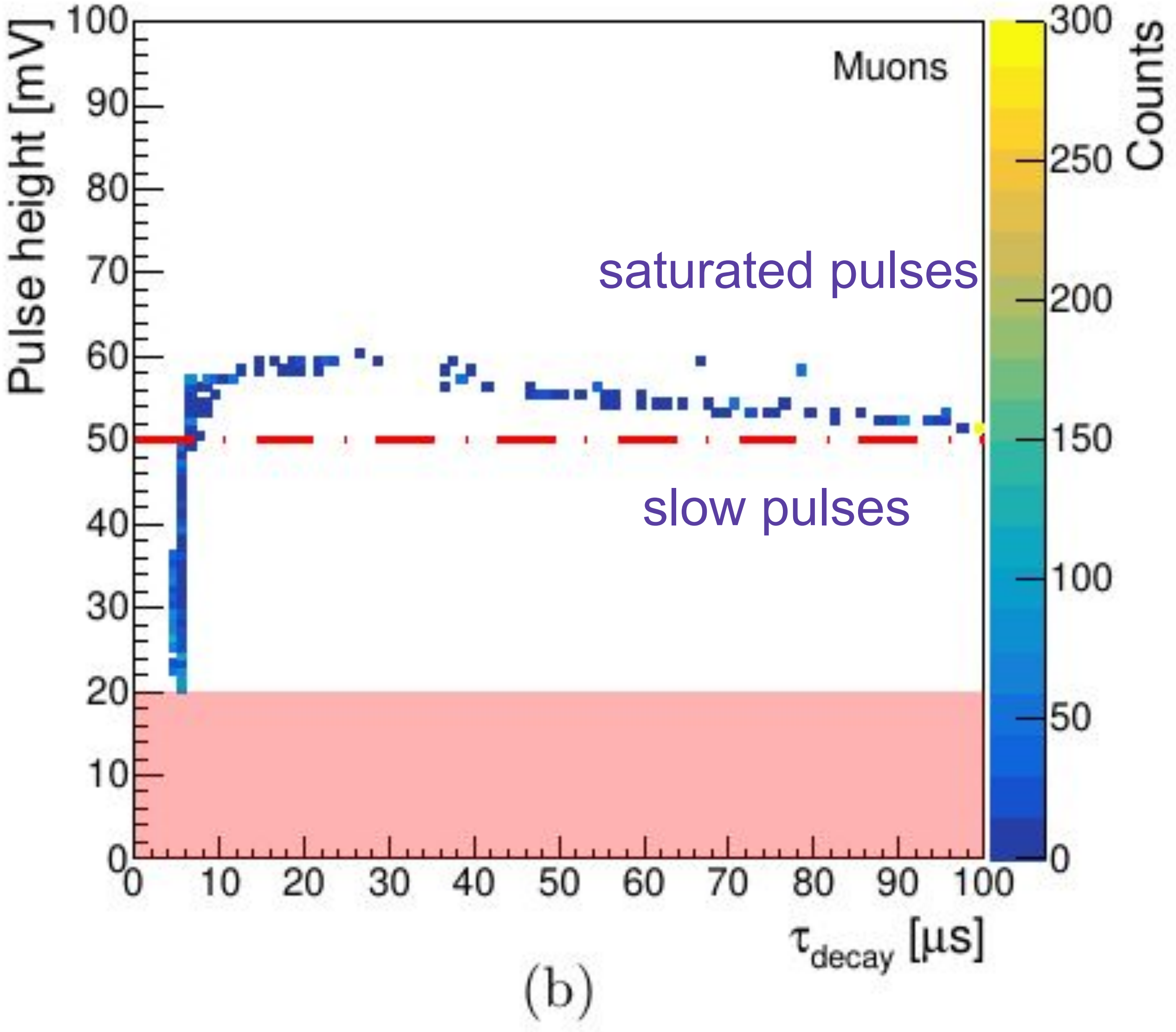
Intrinsic TES Backgrounds



intrinsic background measurement



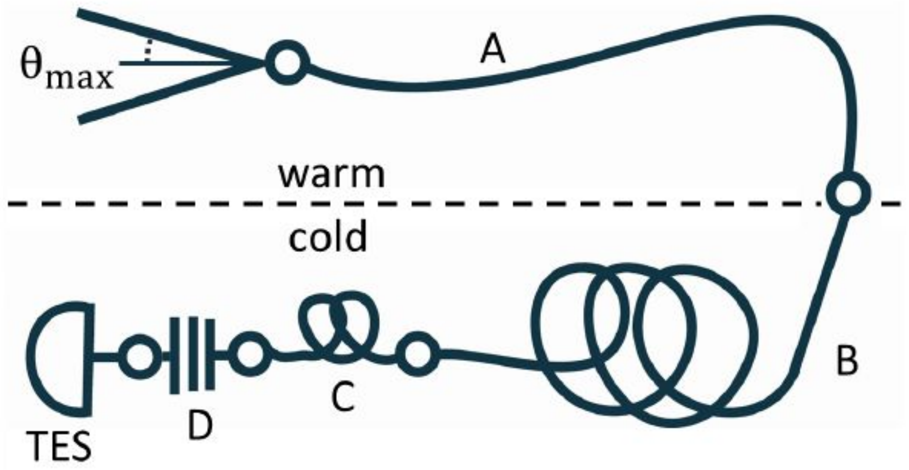
simulation of radioactivity (Zr in fiber sleeve)



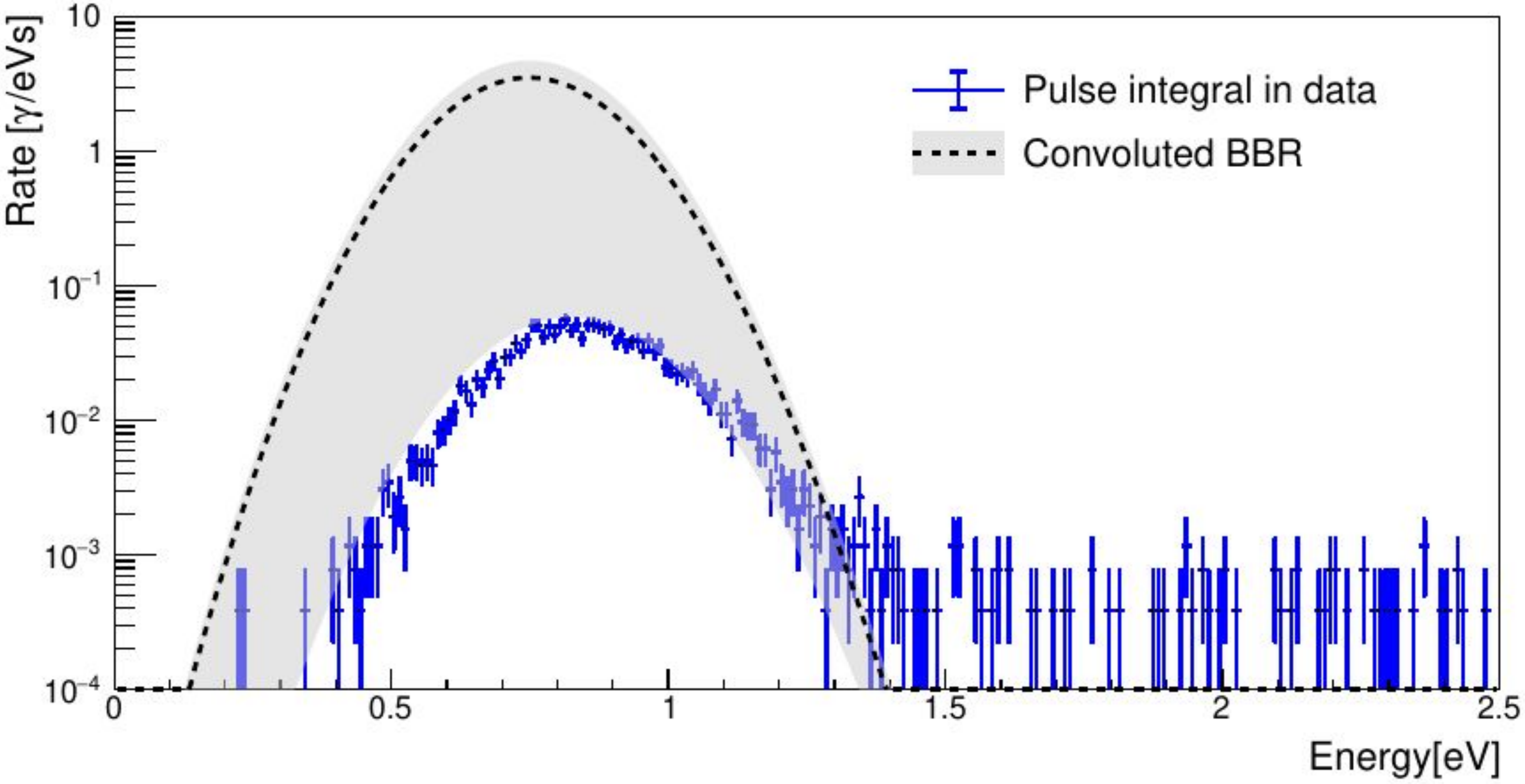
simulation of cosmic muon contribution

Studies by Jose A. Ruberia Gimeno

Extrinsic TES Backgrounds

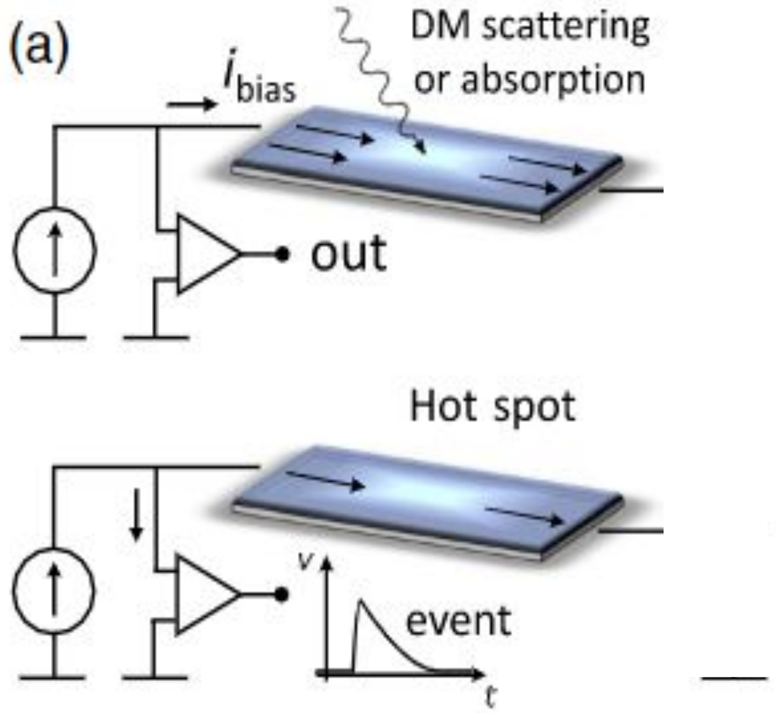


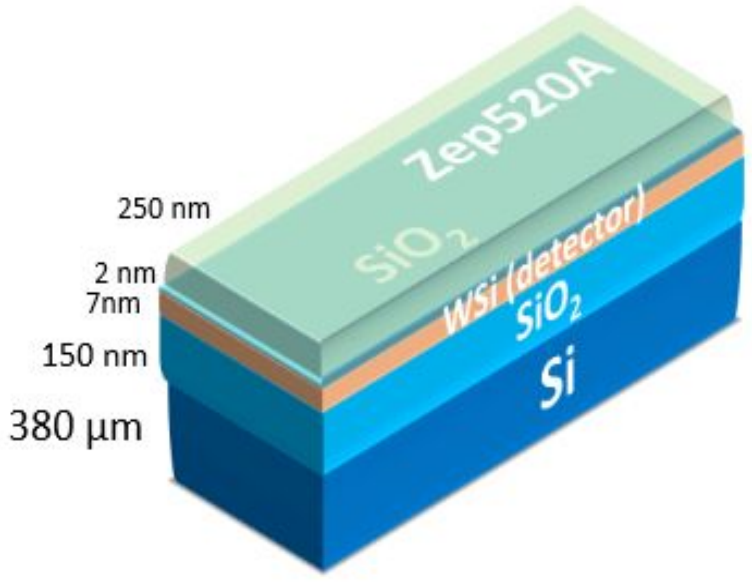
Simulation of black body background contribution considering fiber curling and other components influencing transmission

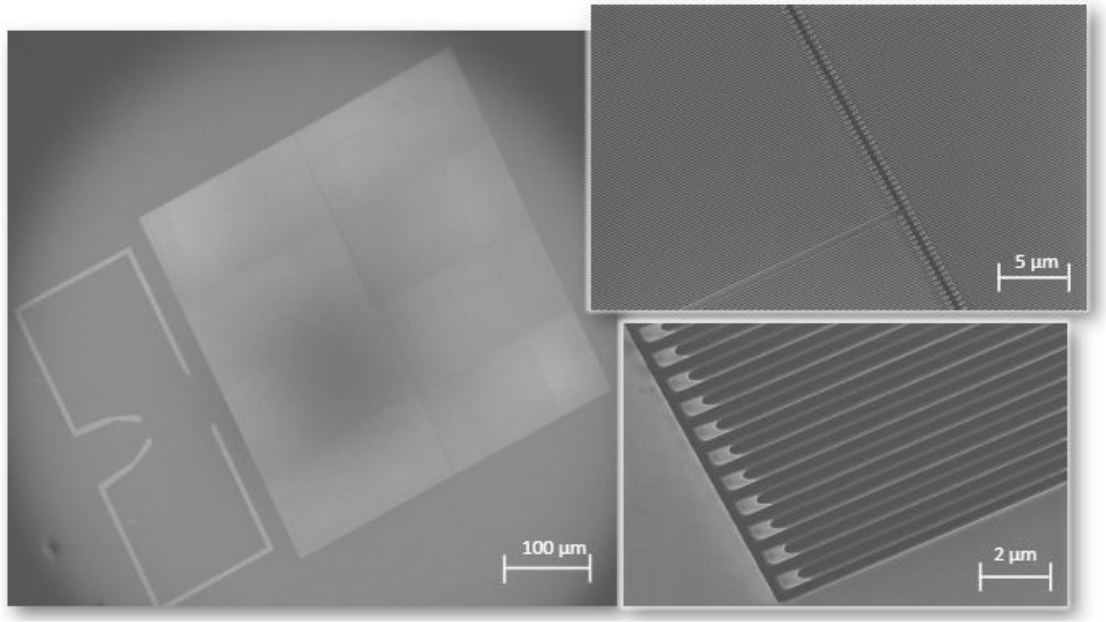


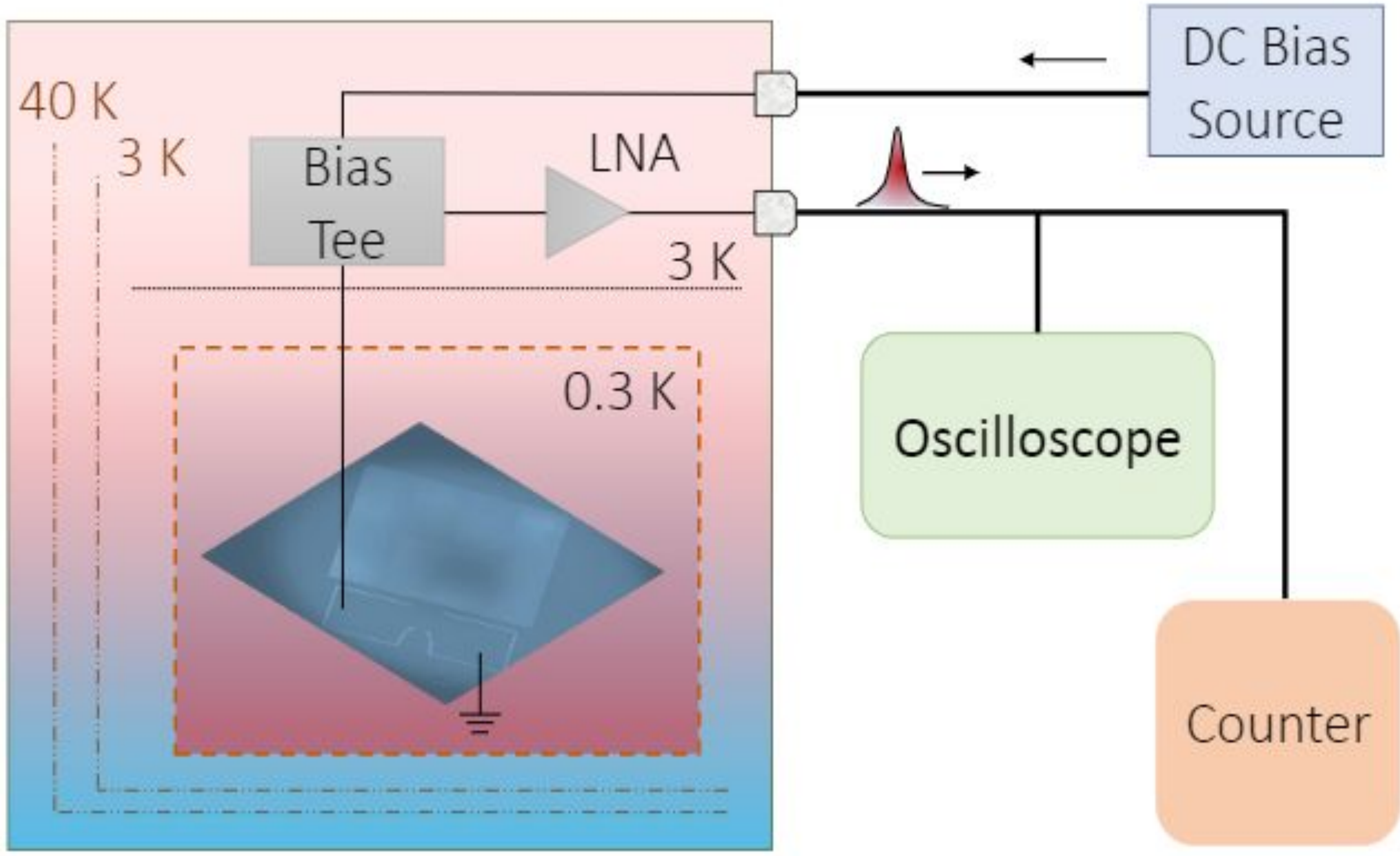
Studies by Jose A. Ruberia Gimeno

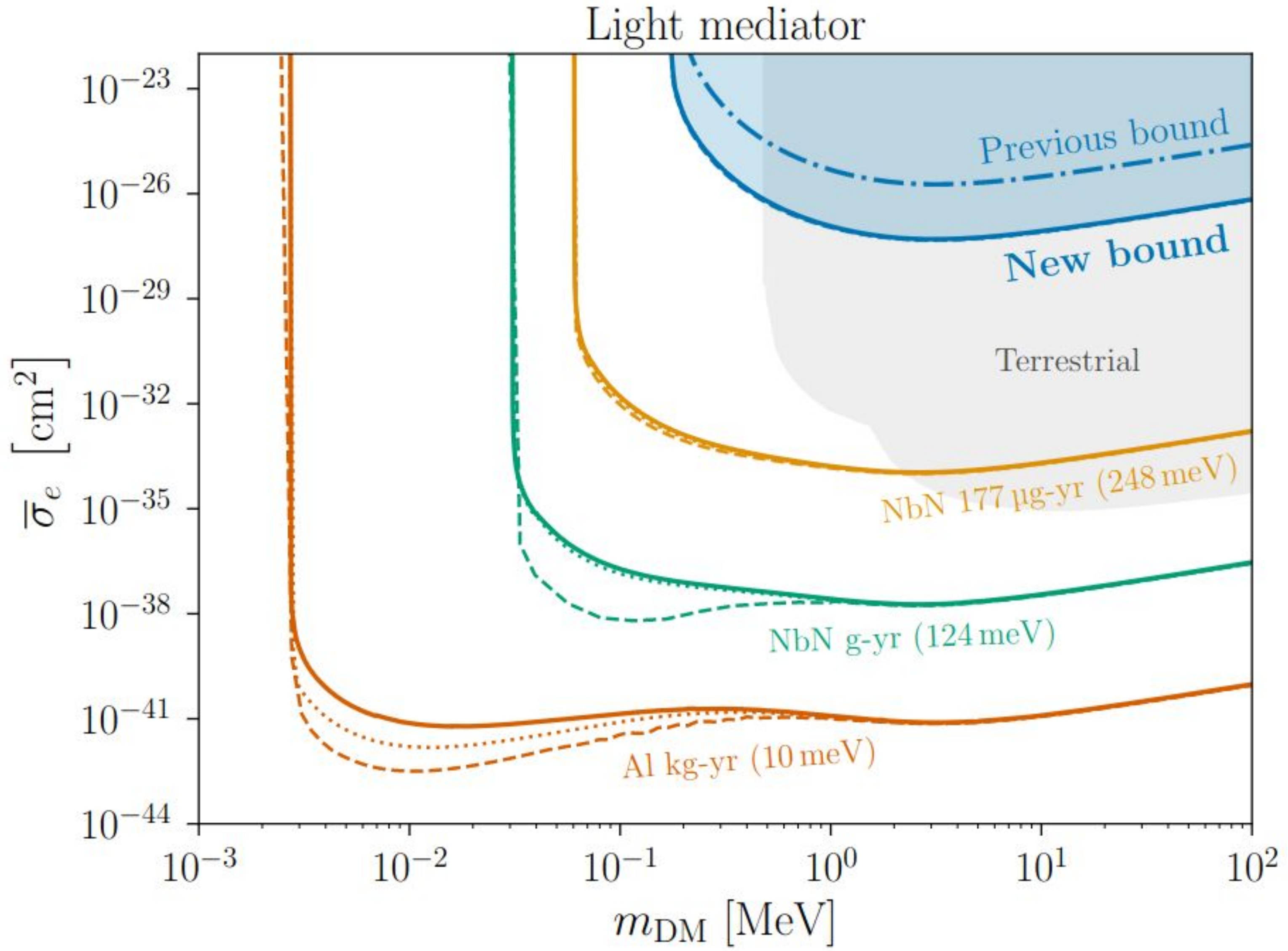
SNSPDs







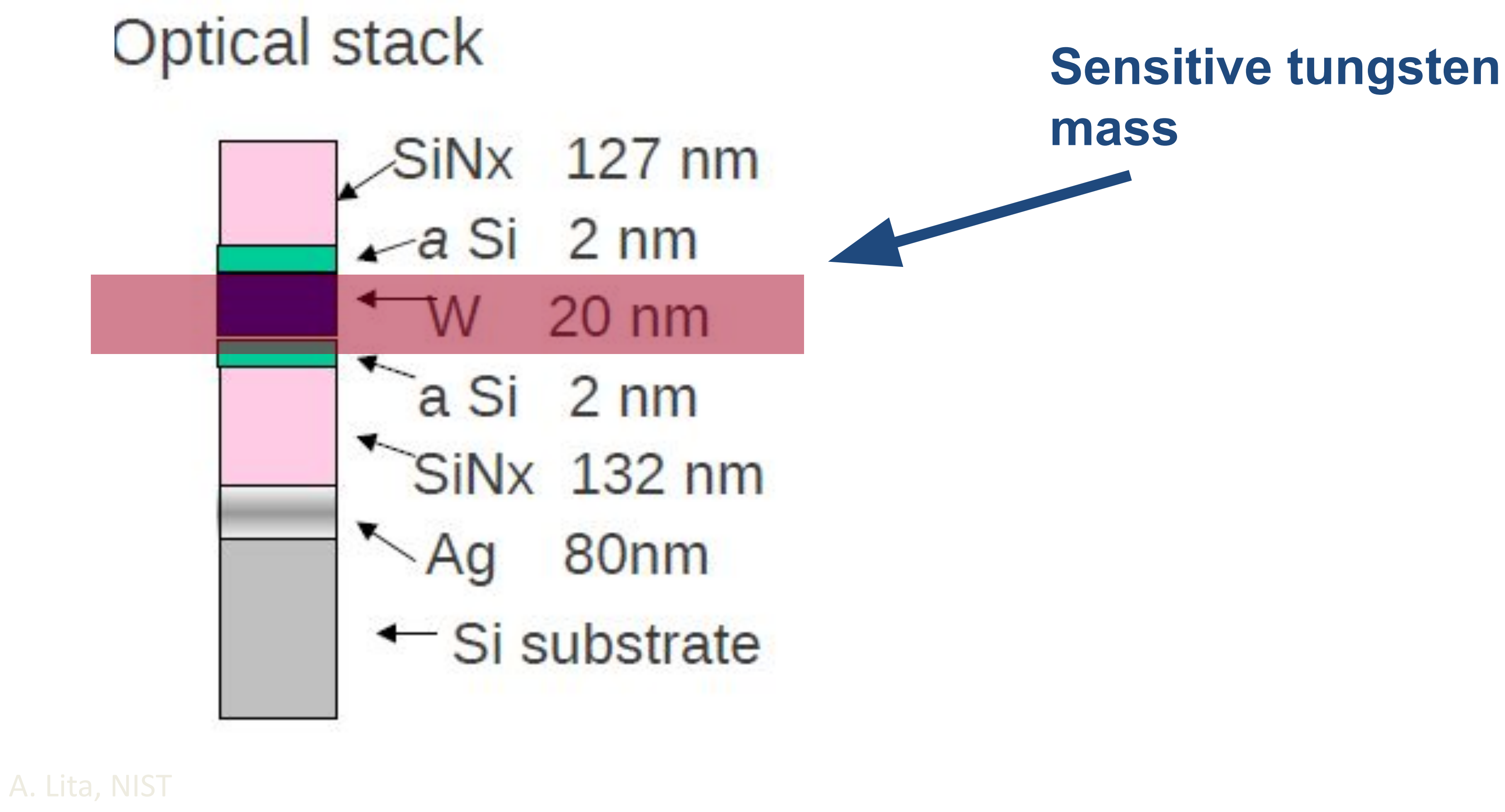


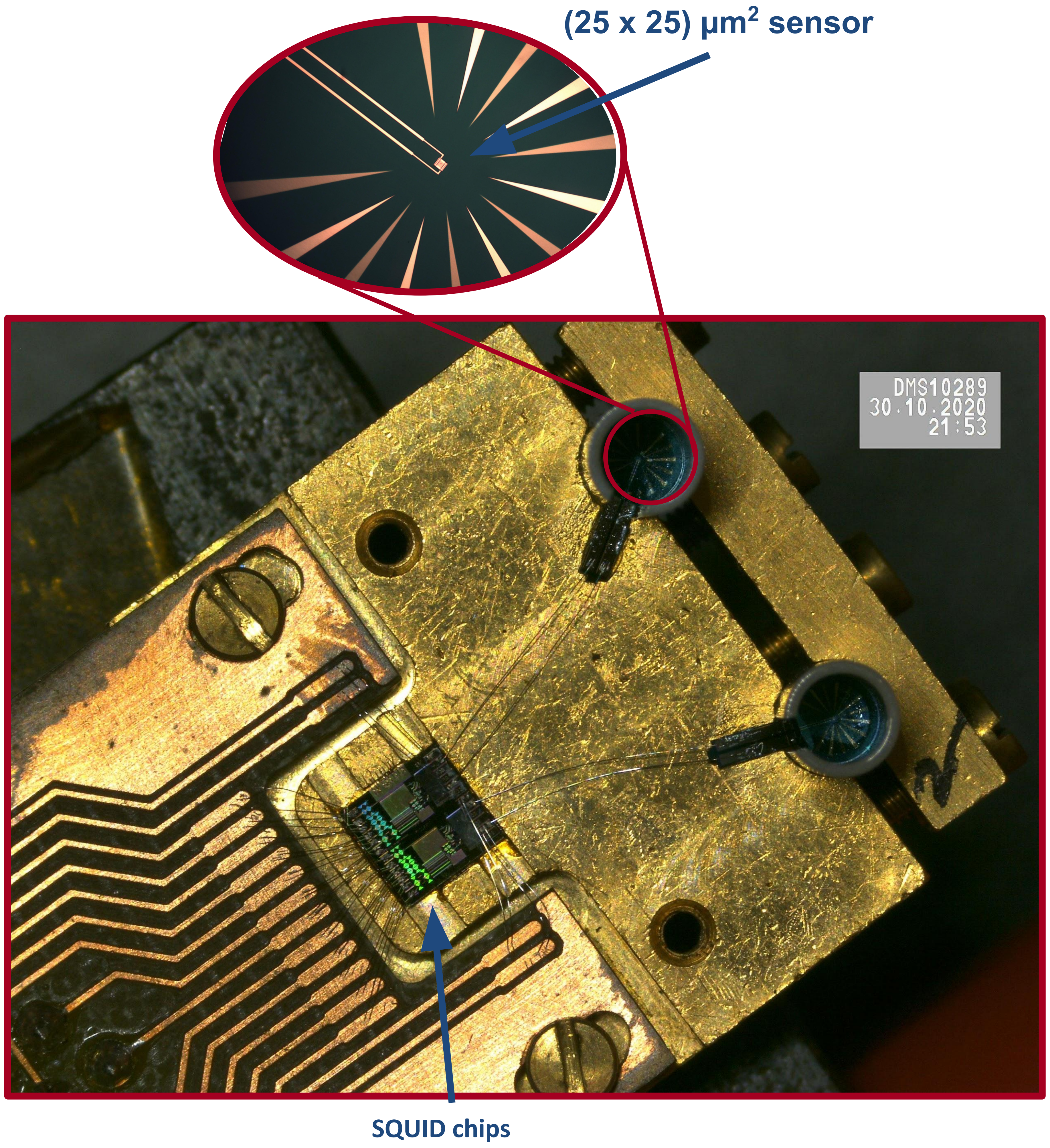


Hochberg, Y. et al. arXiv:2110.01586 (2021)

Hochberg, Y. et al., *Physical Review Letters, 123*(15). (2019)

TES Details





- Very small active area – energy deposition in tungsten layer
- Optical stack & efficiency optimized for 1064nm (1.165 eV) photons
- Wider range of energies interesting for direct DM searches

DDM viability simulations

test background rate for lower triggers after analysis:

- ~ 70 min noise-only simulation
- Applying cuts optimized for 1.165eV and **0.583eV**

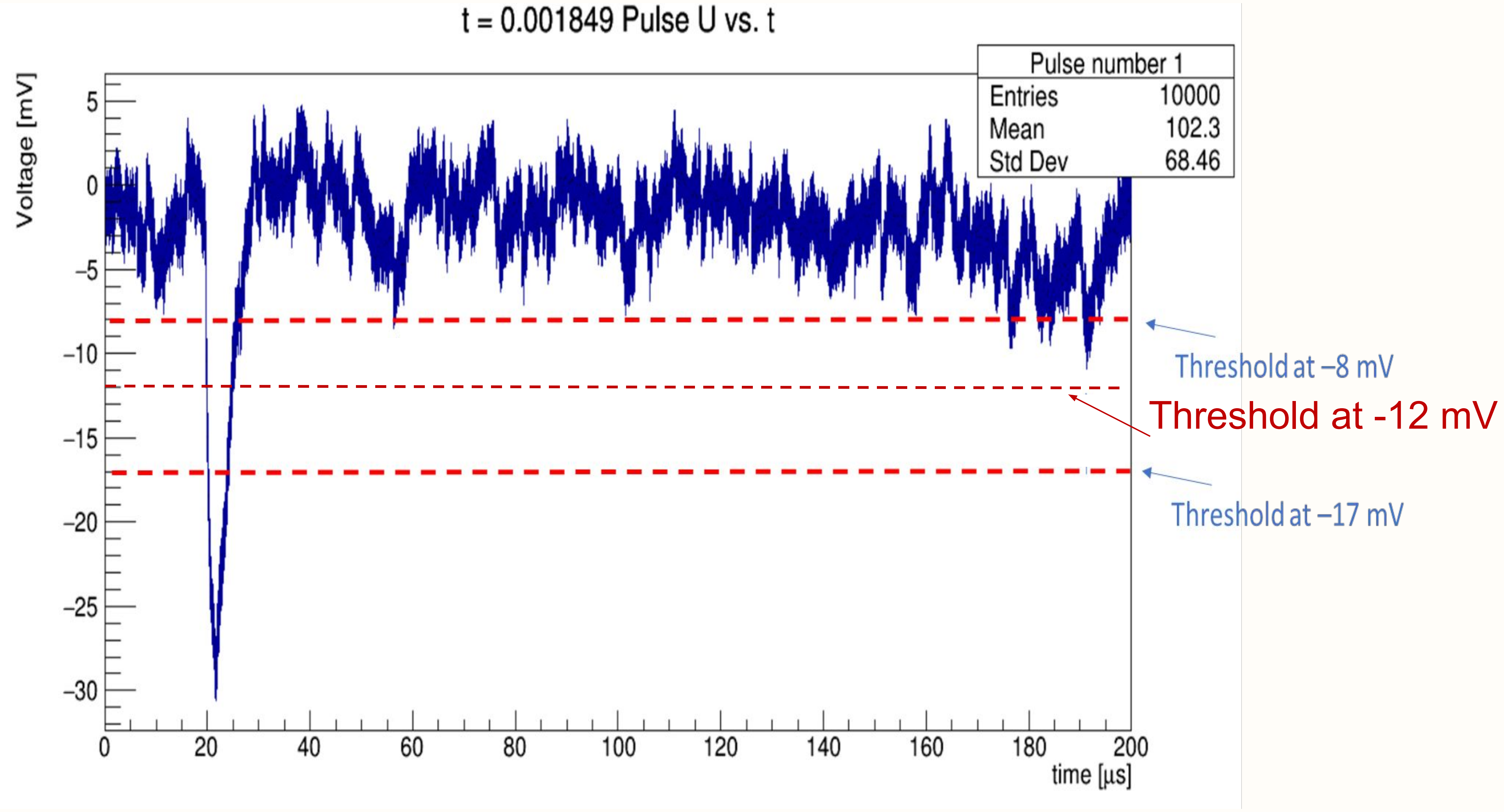
Noise-only simulations

Trigger Rate for -12 mV threshold
0.422 (0.010) Hz

after analysis & cuts

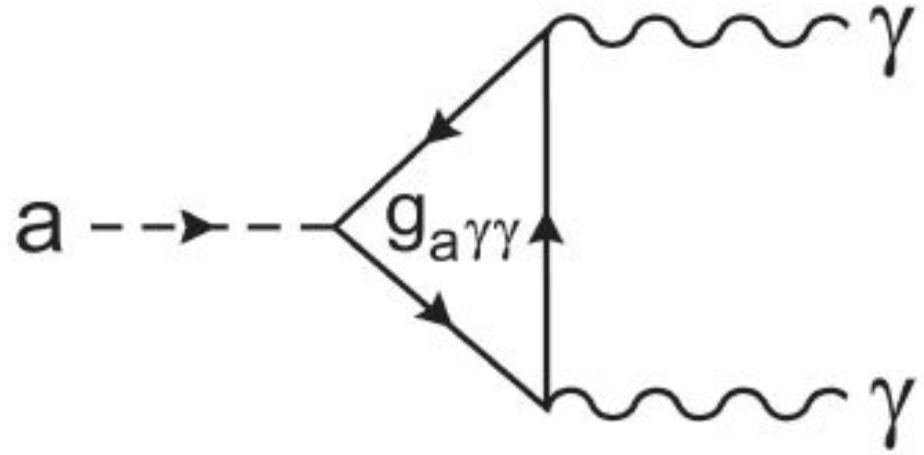
Cuts based on	Trigger Rate for -12 mV threshold
1.165 eV	< 0.0007 Hz
0.583 eV	< 0.0007 Hz

No noise passing analysis & cuts with ~56% acceptance of 0.583eV pulses
Promising for sub-MeV direct DM searches!

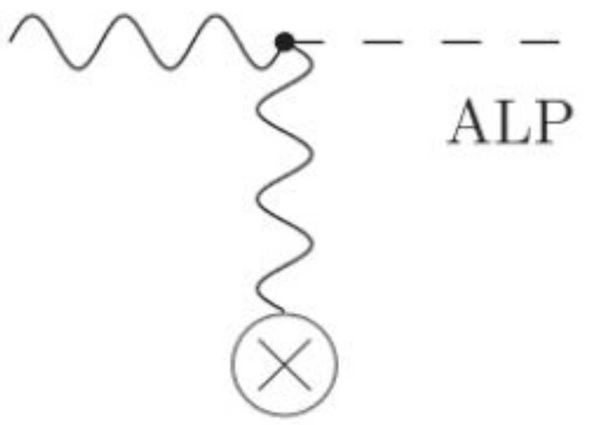


ALPS II - Any Light Particle Search

- SM-coupling to two photons

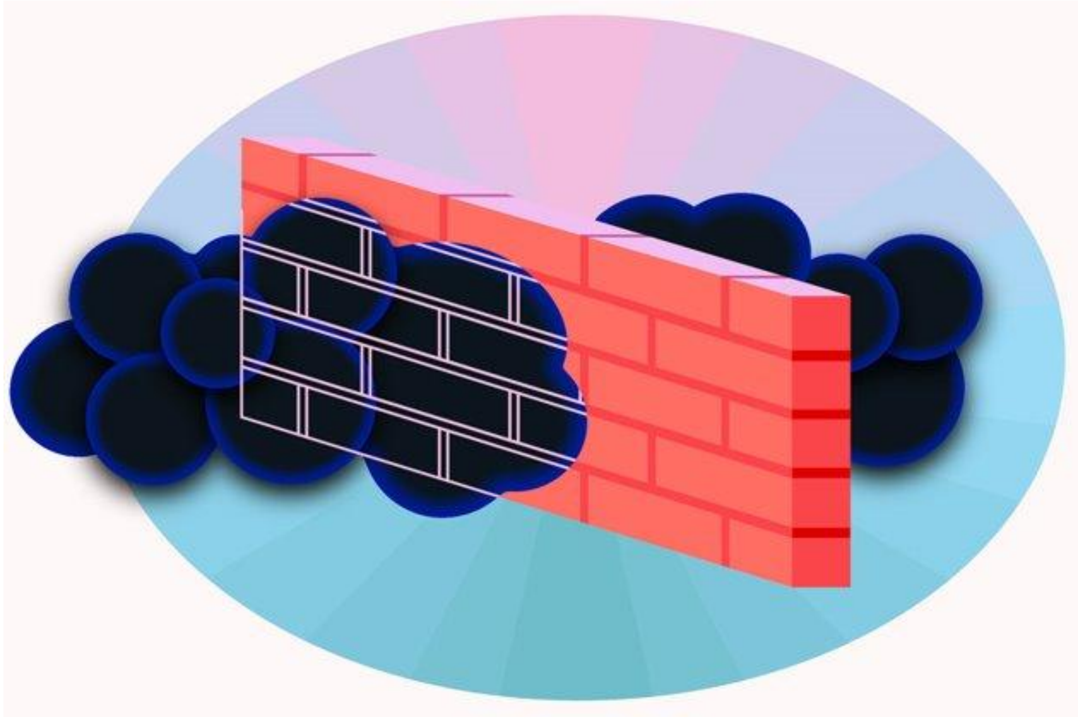


- Detection via Primakoff-like Sikivie effect
- Possible ALP **production** by photon-ALP – oscillation in the presence of strong magnetic fields



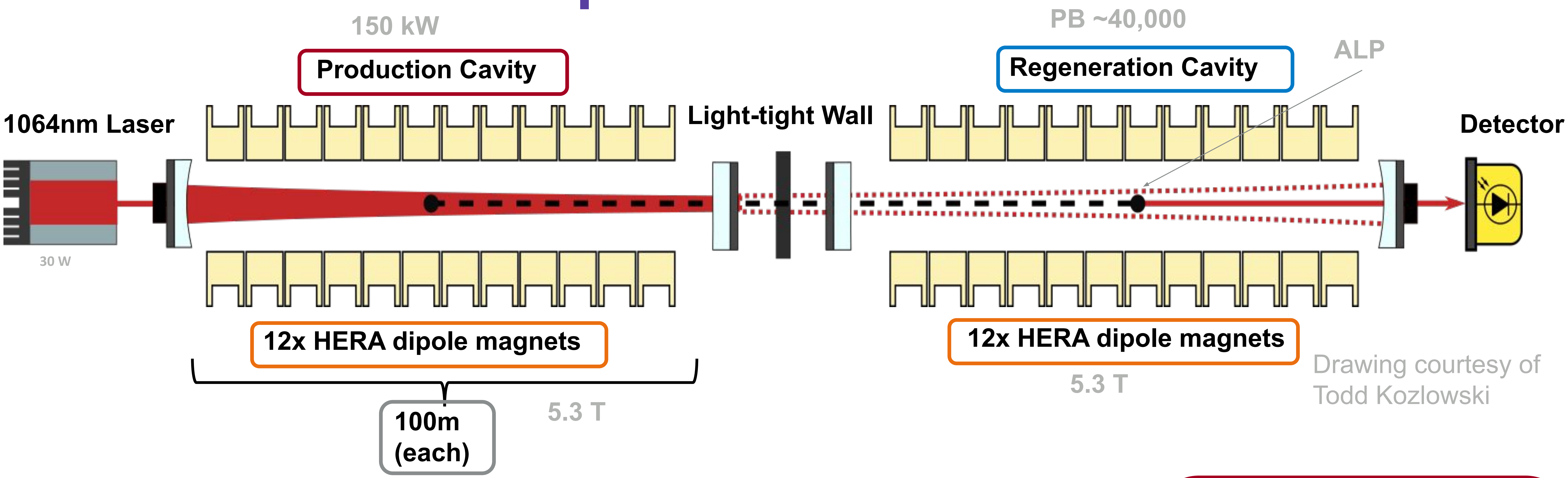
K. Ehret et.al., NIMA 612(1)83-960 (2009)

Light **S**hining Through **W**alls (LSW) experiments



$$\Rightarrow P_{\gamma \to a} \propto g_{a\gamma\gamma}^2 B^2 L^2$$

ALPS II - Setup



Drawing courtesy of Todd Kozlowski

Detection probability:

$$P_{\gamma\to a\to\gamma} \propto PC \cdot RC \cdot g_{a\gamma\gamma}^4 B^4 L^4$$

Expected rate of low energy ($\sim 1.16\,\mathrm{eV}$) photons:
(for $g_{a\gamma\gamma} = 2 \cdot 10^{-11}\,\mathrm{GeV}^{-1}$)

$$\dot{N}_\gamma \approx 2.8 \cdot 10^{-5} \frac{\gamma}{\mathrm{s}} \approx 1 \frac{\gamma}{\mathrm{day}}$$

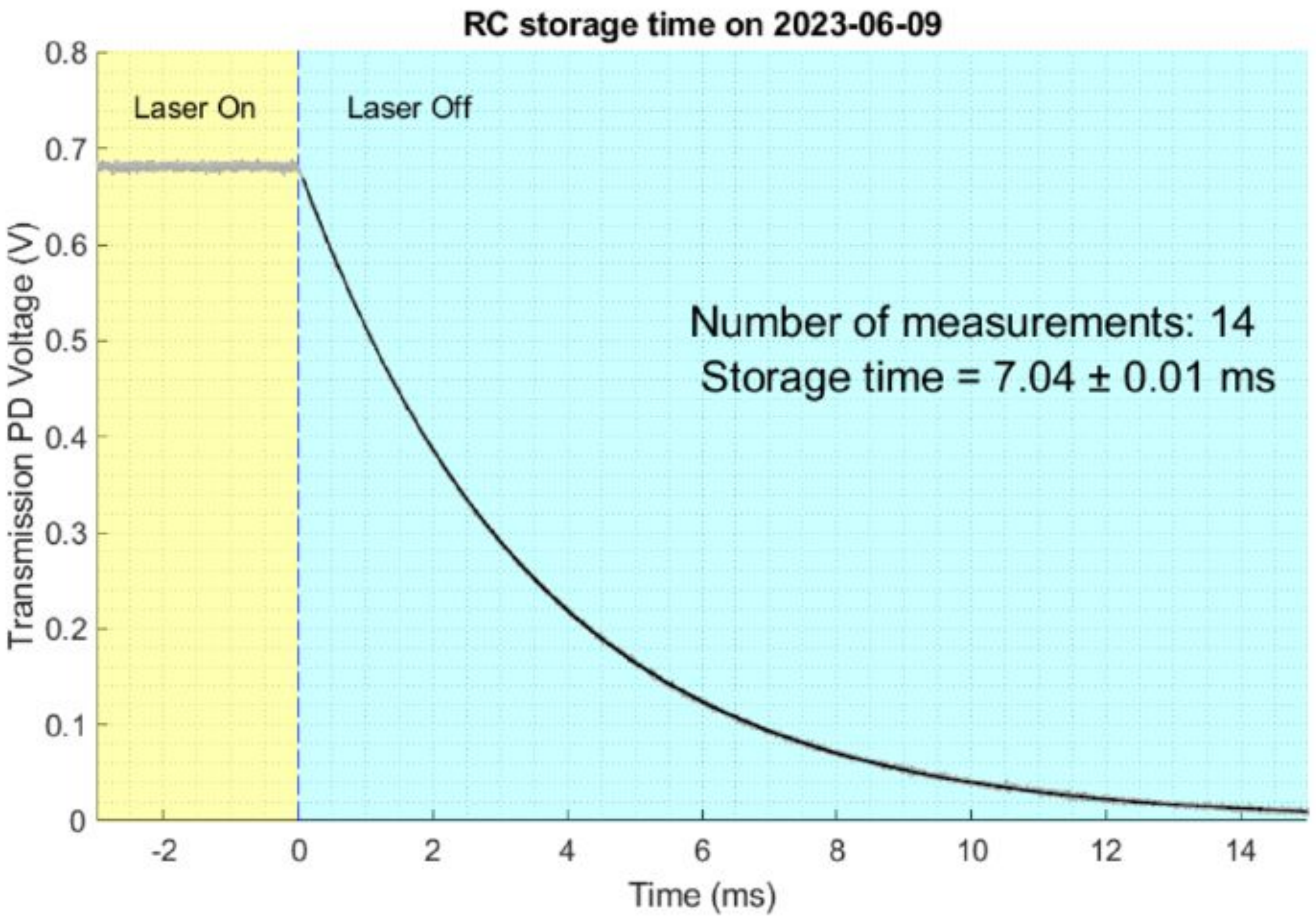
Single-photon detection requirements for ALPS II:

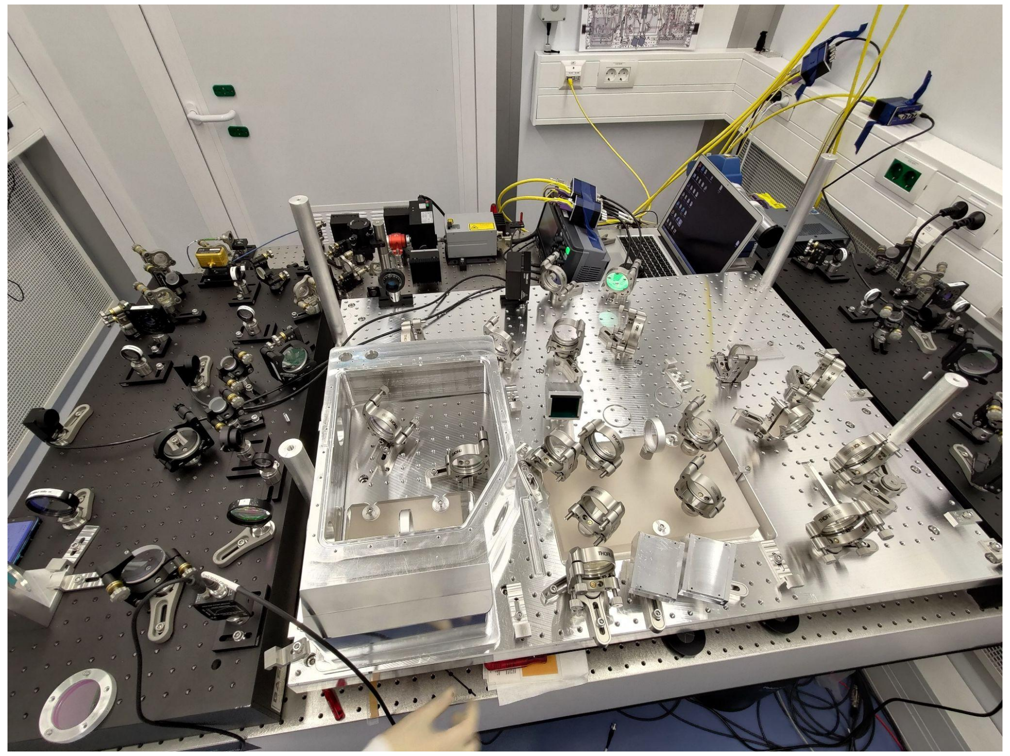
- Low energy photon detection
- Low background (< 1 photon/day)
- High detection efficiency

ALPS II - World leading precision interferometry

- Longest storage time Fabry Perot cavity ever!
- Length: 124.6m, FSR: 1.22 MHz
- Storage time: **7.04 ms**

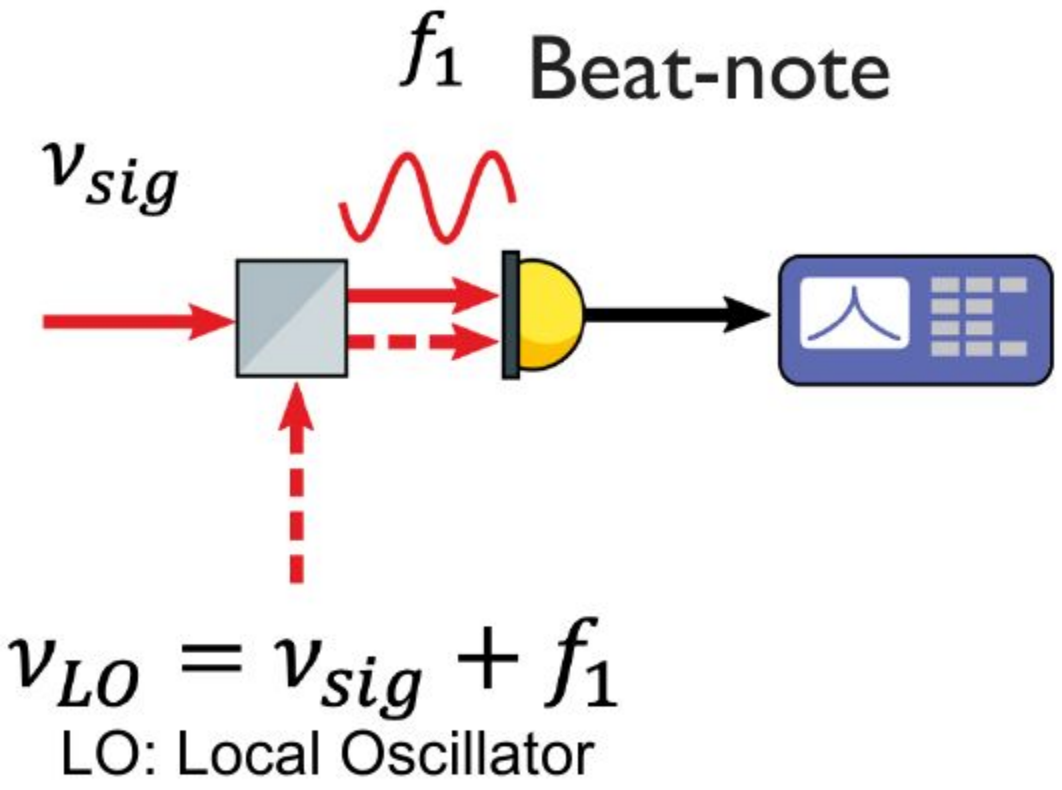
Leading precision interferometry!

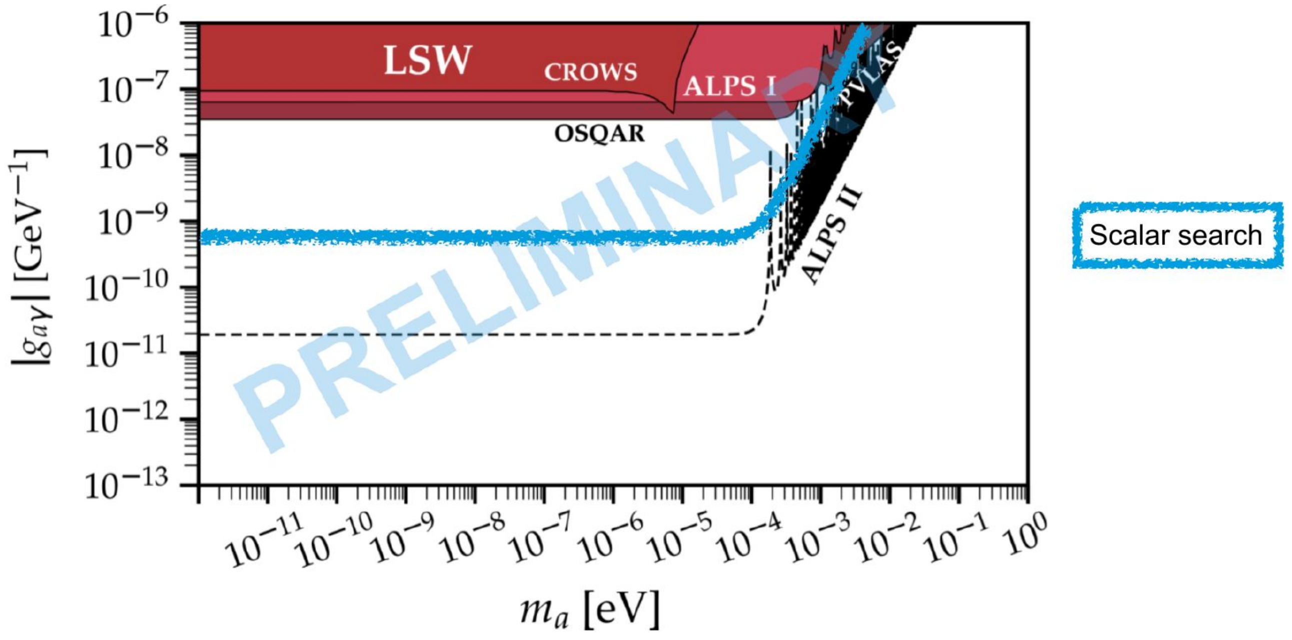


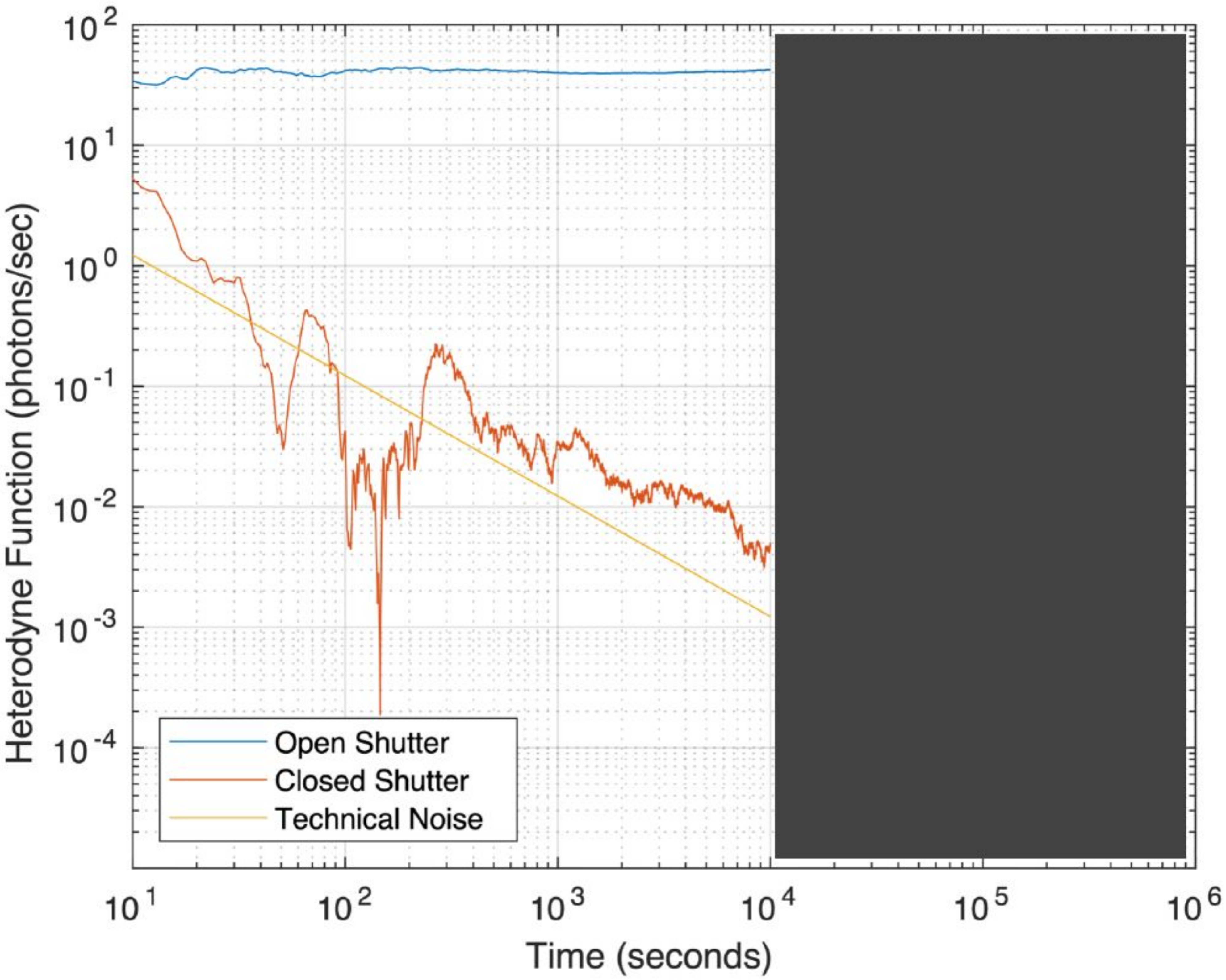


Slides adapted from Isabella Oceano

ALPS II - Heterodyne detection - first Results (Production Cavity only)

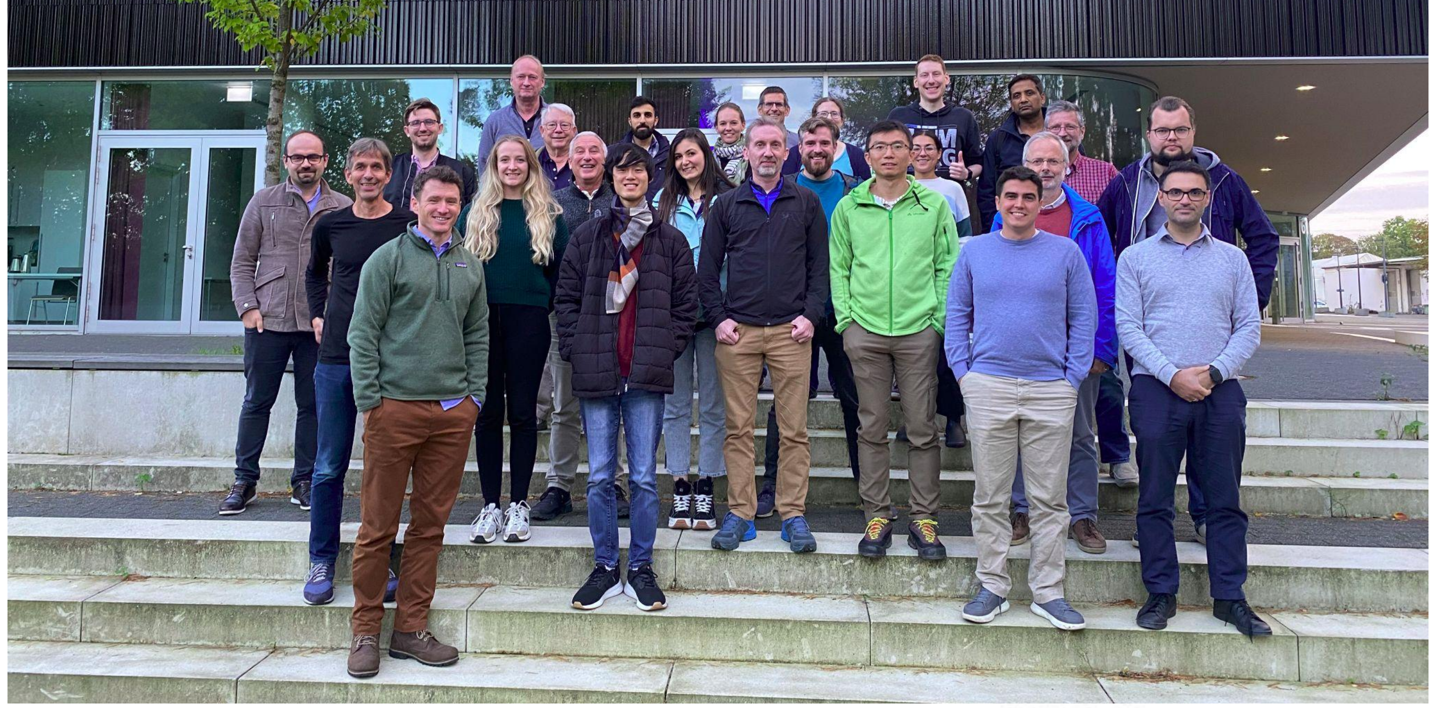






Slides adapted from Isabella Oceano

ALPS II - Timeline



2023 | 2024 | 2025 | 2026

May | January | April | July ... | ...

- Engineering run
  - ~150k s integration time
- Fix phase problems
- Improving the SNR
  - ~2.5 factor
- Automatise the locking procedure
- Scalar search
  - ~800k s integration time
- Pseudoscalar search
  - 1M s integration time
- Installation of the production cavity
- Pseudoscalar search
- Design sensitivity

Slides adapted from Isabella Oceano