Gravitational Waves from the innermost parts of Core-Collapse Supernovae

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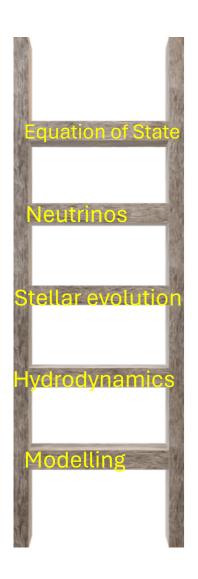
By Maria Pereira 💠











Since 1932, more than 160.000 papers have been written (key words "Supernova" or "Neutron star")

Since 2014 ~7000 Supernovae discovered

SN1987A

1985 Bethe & Wilson

1966 Colgate & White

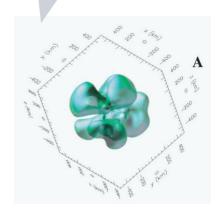
1934 Zwicky & Baade





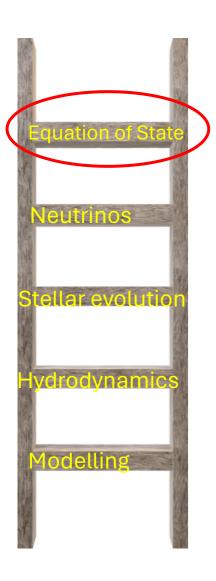
Cover of TIME magazine

Early 2000s (first 3D simulation)



3D by Fryer & Warren (2002)

In addition, the new problem of developing a more detailed picture of the happenings in a super-nova now confronts us. With all reserve we advance the view that a super-nova represents the transition of an ordinary star into a neutron star, consisting mainly of neutrons. Such a star may possess a very small radius and an extremely high density. As neutrons can be packed much more closely than ordinary nuclei and electrons, the



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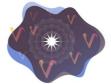
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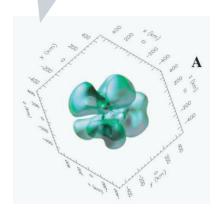






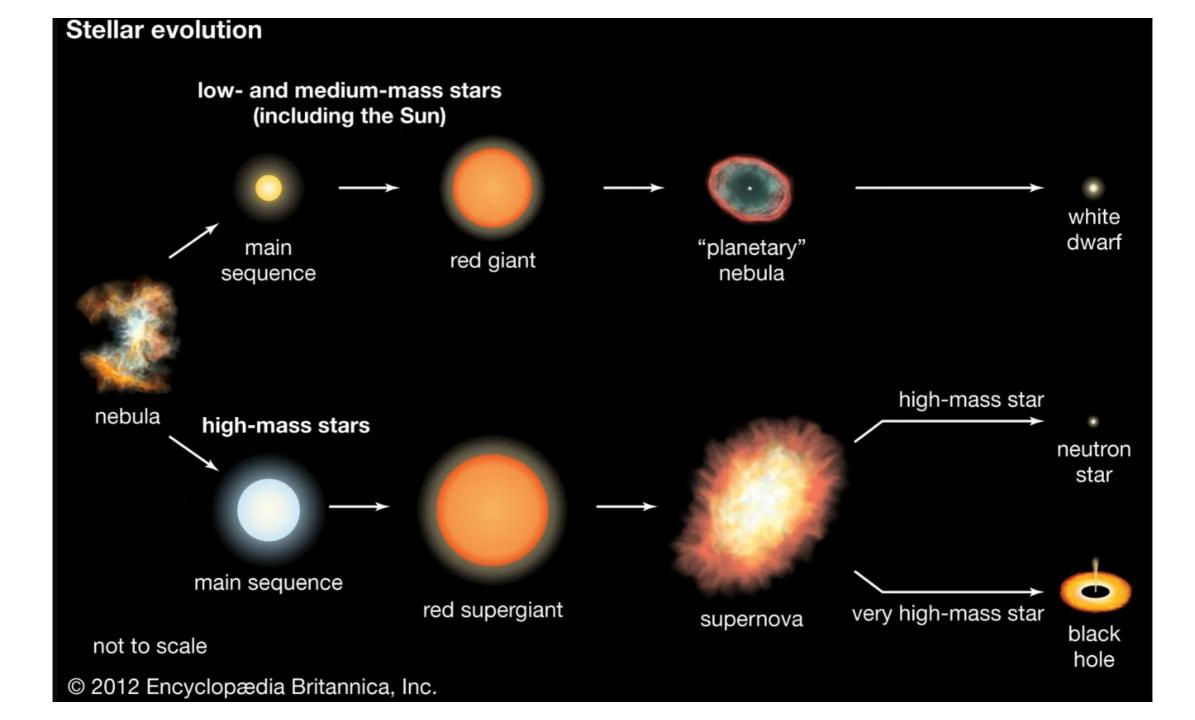
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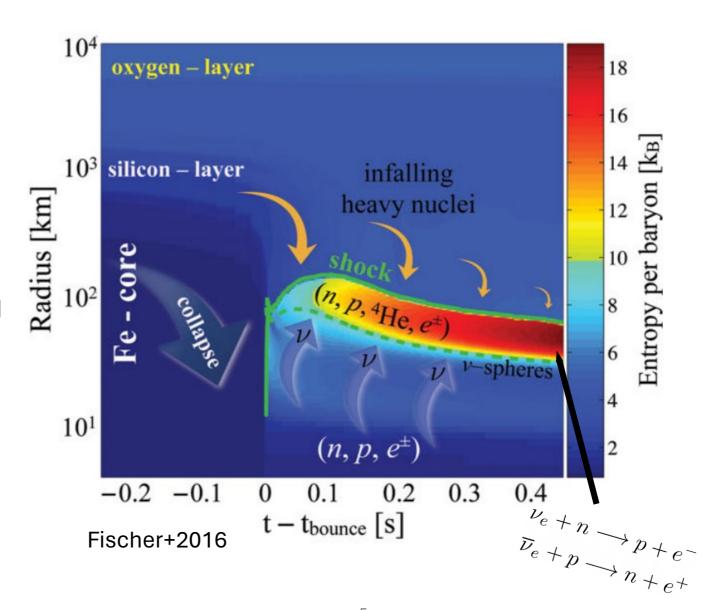
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I: Collapse phase

- •Iron core: No exotherm reaction possible
- Iron core accreting material from Silicon shell burning
- Exceeds Chandrasekhar mass limit
- "Time zero" of what (can) be called a CCSN
- Core collapses

Collapse further accelerated by electron captures on Iron peak nuclei



II: Core-bounce

Around <u>nuclear saturation</u> <u>density</u>:

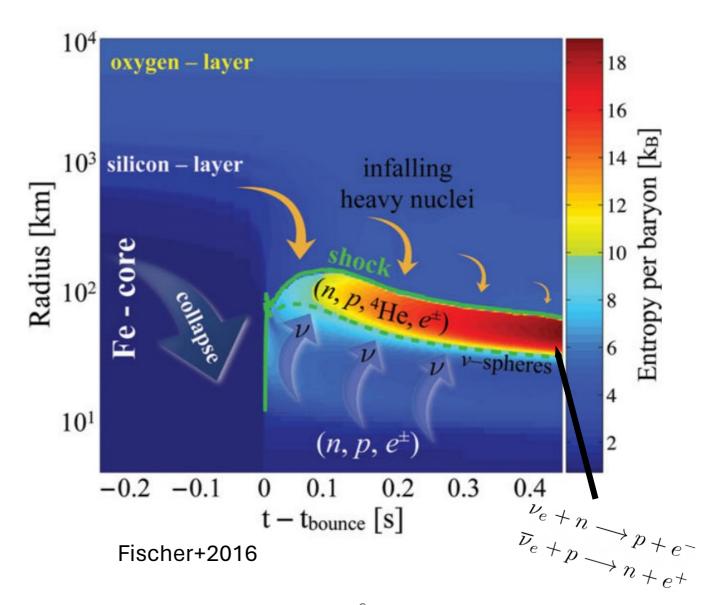
Neutrinos become trapped

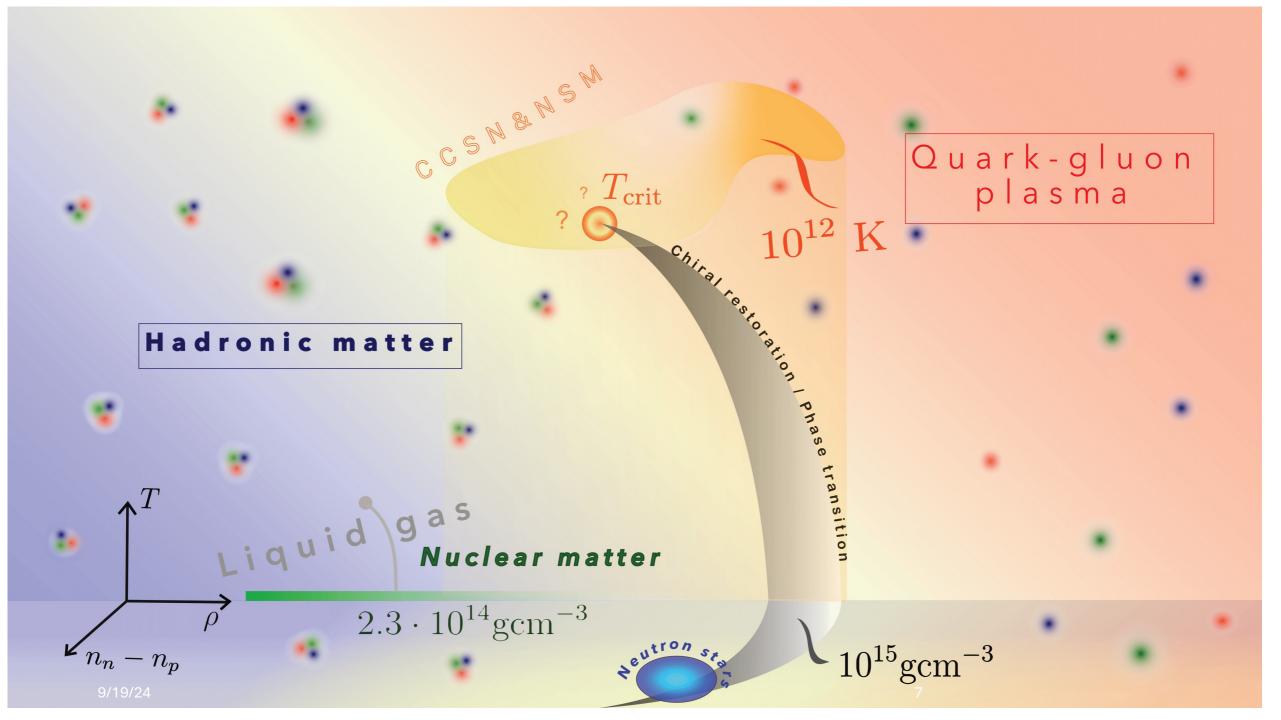
- → Iron core can either form a BH or a NS
- → Core decelerates and bounces
- → Shock wave forms ~1e51 erg

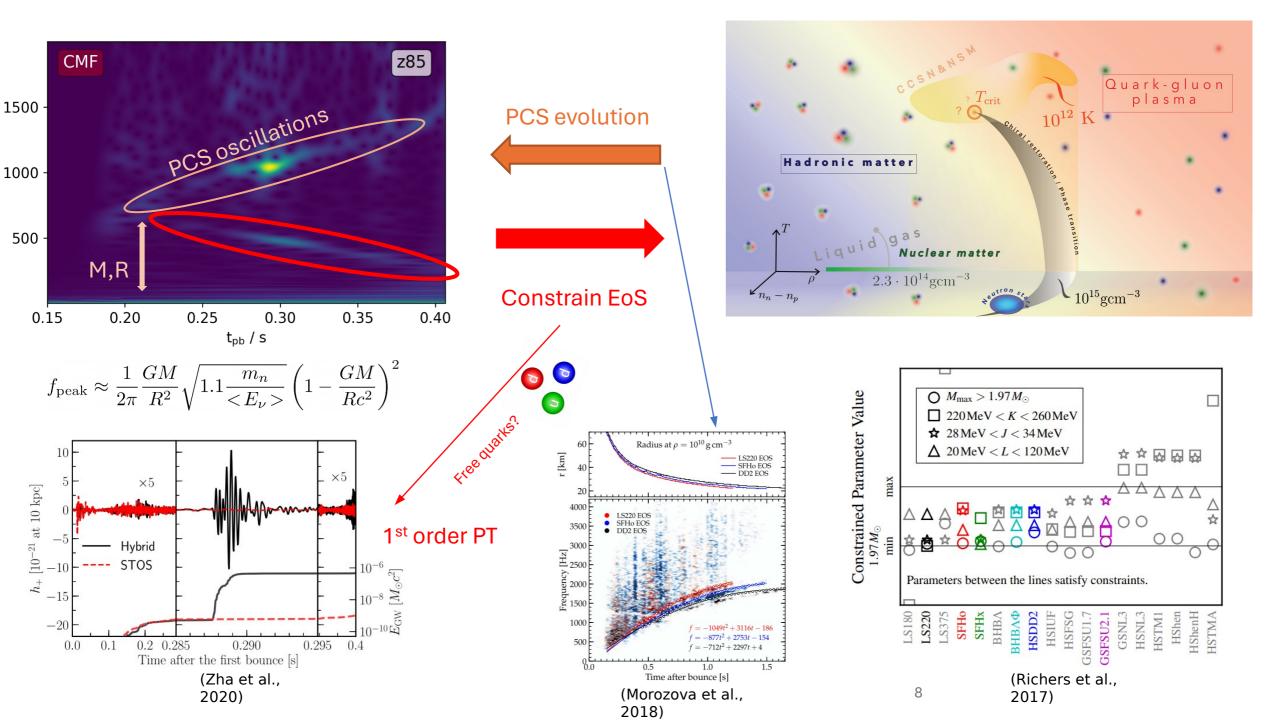
III: Explosion phase

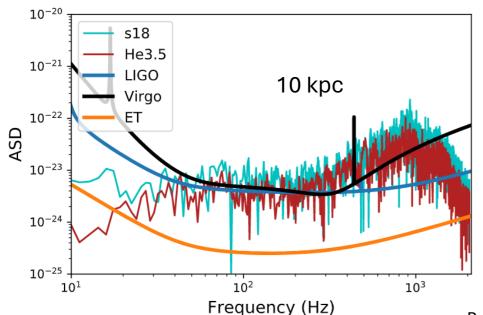
Shock wave stalls...

- "real explosion" sets in 100-200ms post-bounce
- Driven by neutrino heating behind the shock
- Explosion can outshine entire galaxy!
- Asymmetric effects pivotal





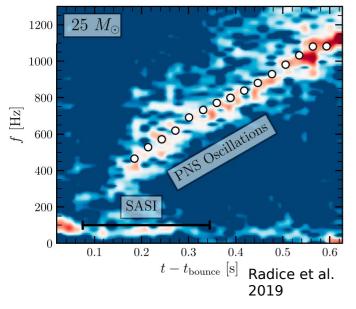




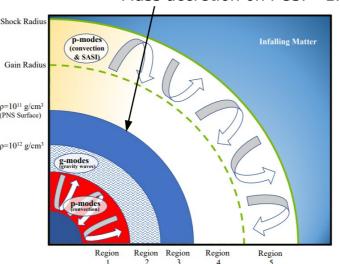
Gravitational Waves from Core-Collapse Supernovae

- CCSNe ~3 event per century in our Galaxy
- Observable distance ~10-20 kpc
- Time changing mass quadrupole moment
- GWs provide access to compact inner core

Powell & Mueller (2018)



Radius of Proto-Compact Stars \sim 40km to \sim 10km Mass accretion on PCS: \sim 1.4/2



Mezzacappa et al. 2020

- Different regions susceptible to different mode oscillations (p-,g- f-modes)
- Positive entropy gradient: g-mode oscillations
- Characteristic frequency of g-modes:

Print Väicälä froguonov

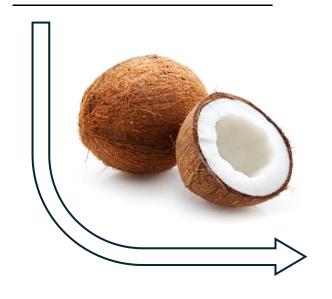
$$\omega_{\rm BV}^2 = \frac{\mathrm{d}\alpha}{\mathrm{d}r} \frac{\alpha}{\rho h \Phi^4} \cdot \left(\frac{1}{c_s^2} \frac{\mathrm{d}P}{\mathrm{d}r} - \frac{\mathrm{d}\rho}{\mathrm{d}r} \right)$$

Stripy clouds



Setup

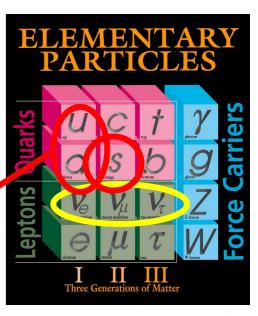
CoCoNut-FMT







Incorporated hadron-quark EoS into CoCoNuT-FMT



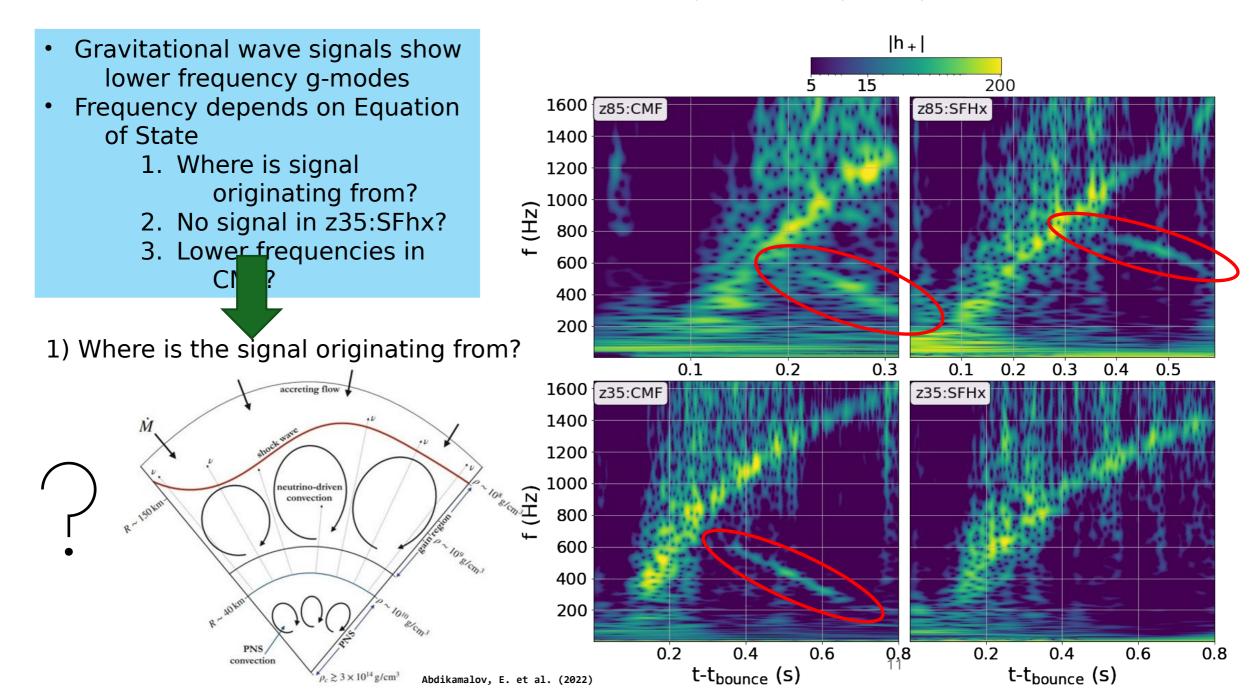
neutrinos.fnal.gov

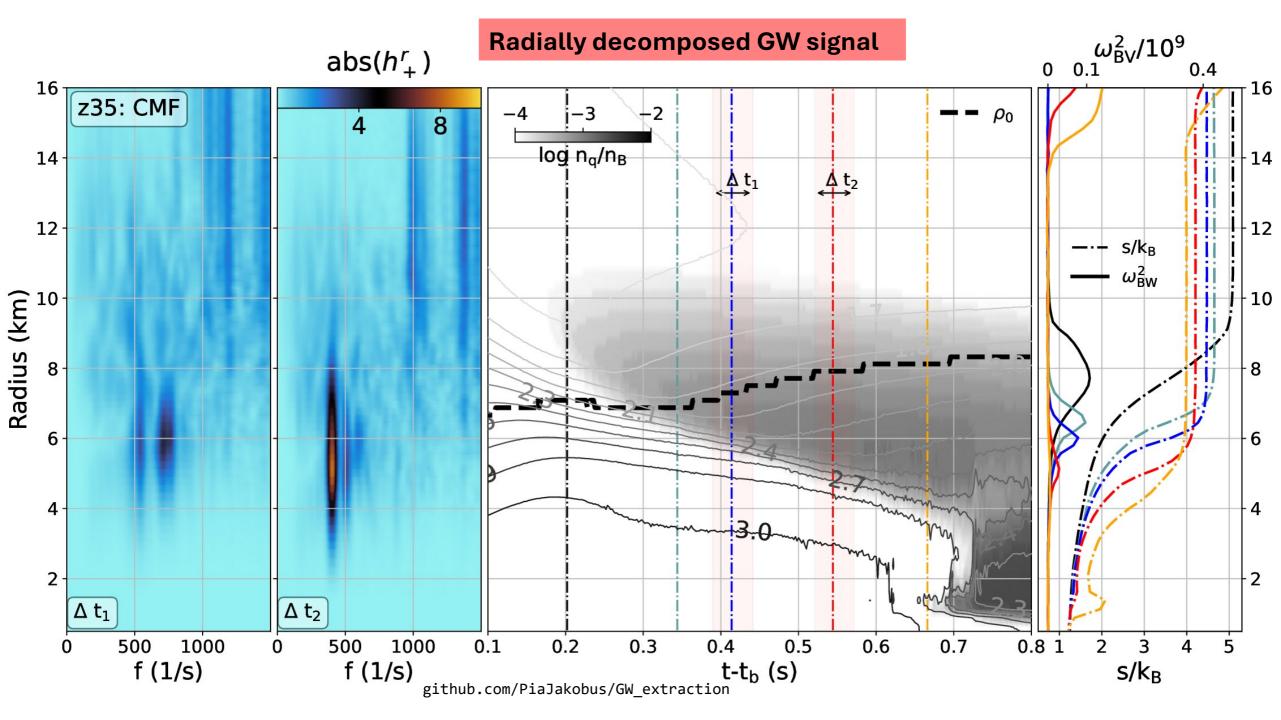
- ➤ **Godunov-based** finite-volume solver (Dimmelmeier et al. 2002; Müller et al. 2013)
- > 3+1 splitting formalism; CFC approximation
- Three flavour Fast Multigroup Transport (FMT)

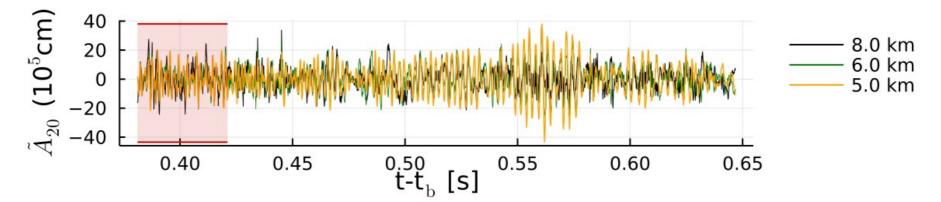
Up to 50 million core hours ©

CoCoNuT-FMT: Müller et al. (2010), Müller & Janka (2015) Kepler: Weaver et al. (1978), Heger & Woosley (2010)

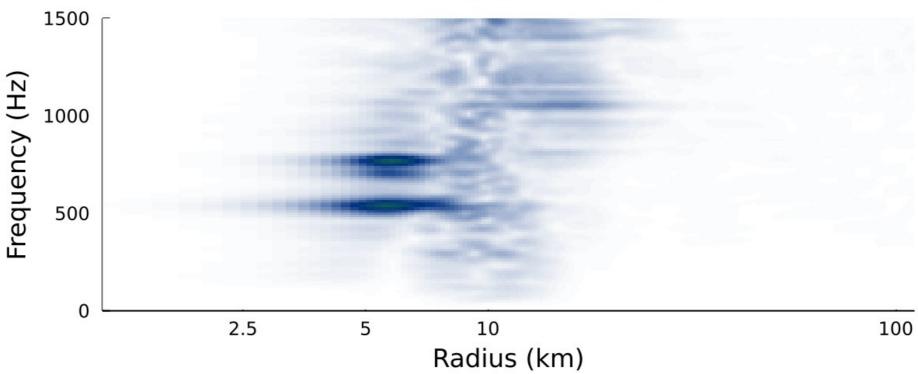


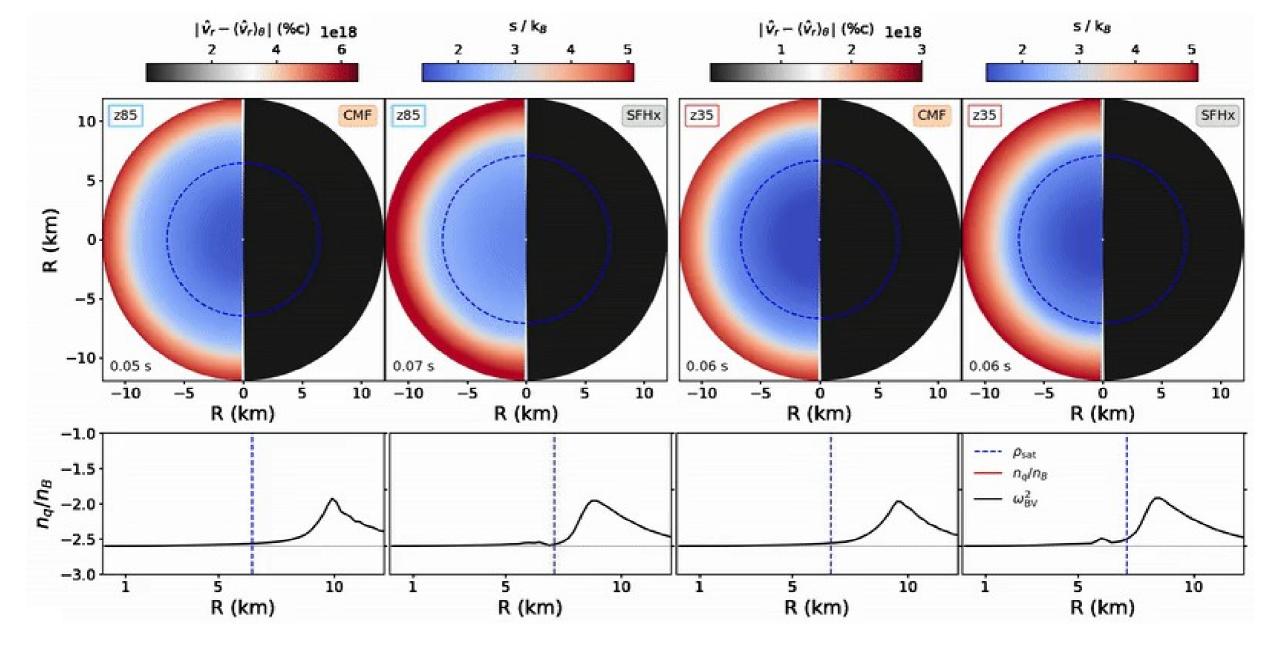


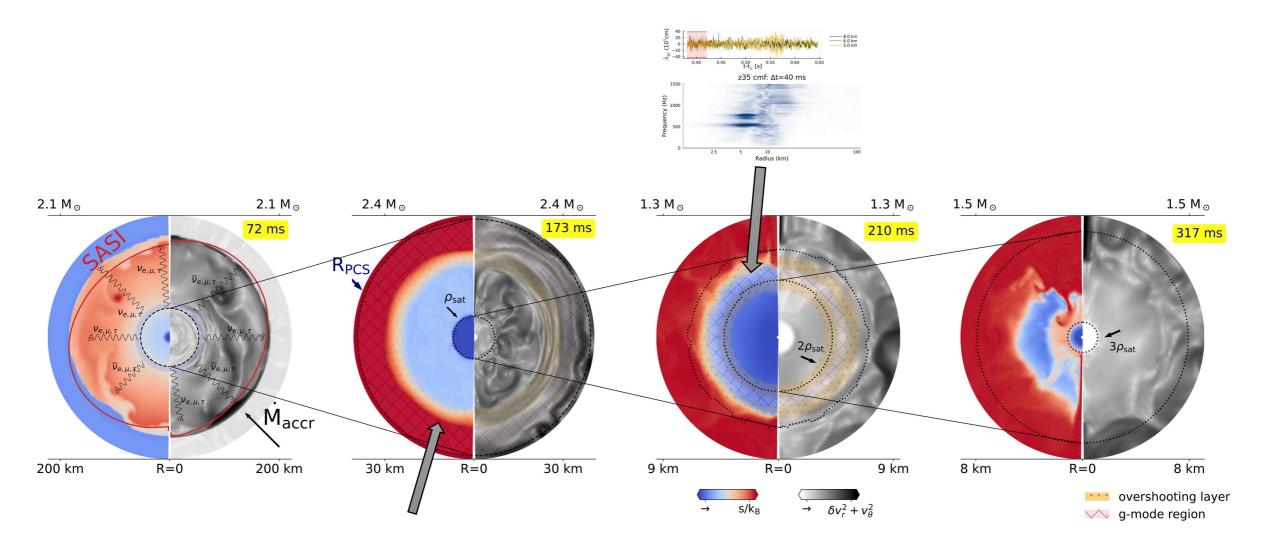


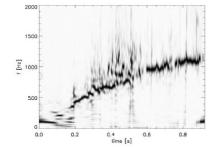


z35 cmf: Δt =40 ms



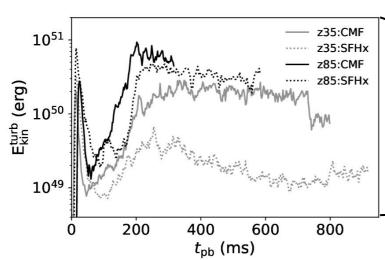




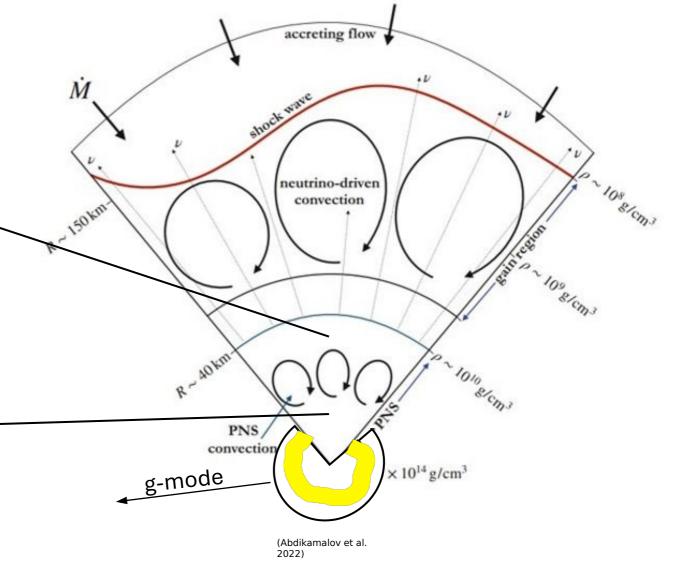


Where is signal originating from?

- 1. g-mode "lives" in convectively stable PCS region beneath the PCS convection zone
- 2. Higher turbulent convective energies seen in CMF, particularly z35





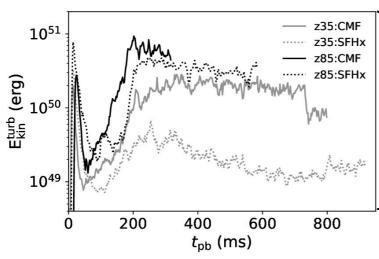


No signal in z35:SFhx?

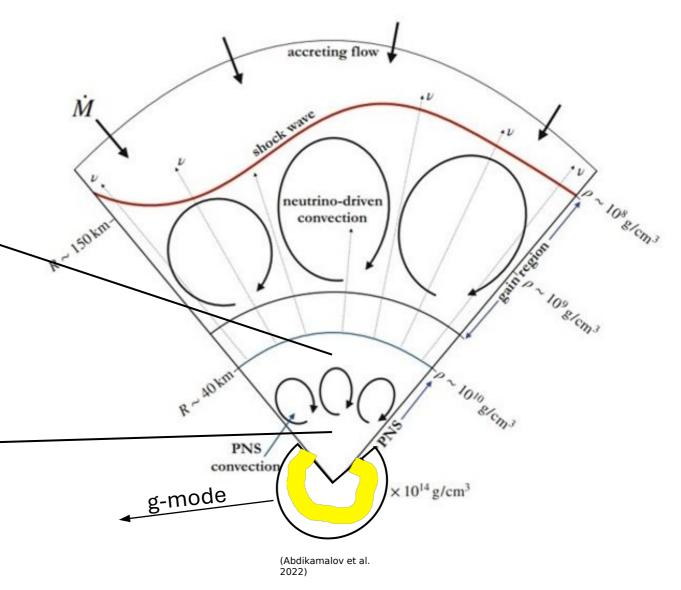
1.

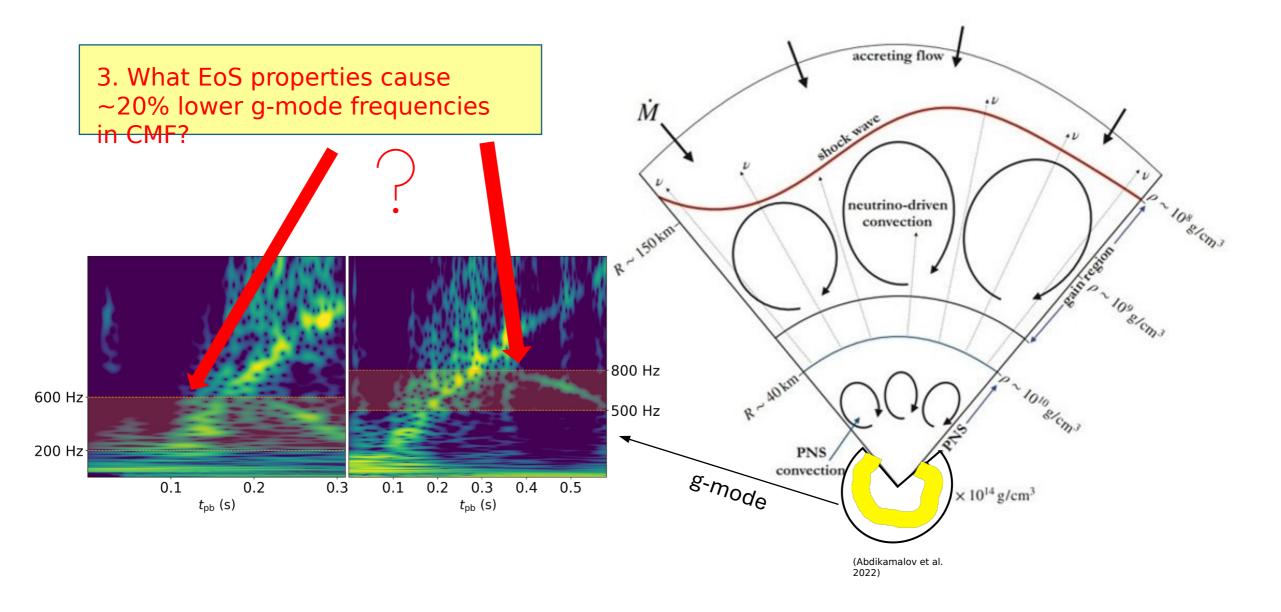
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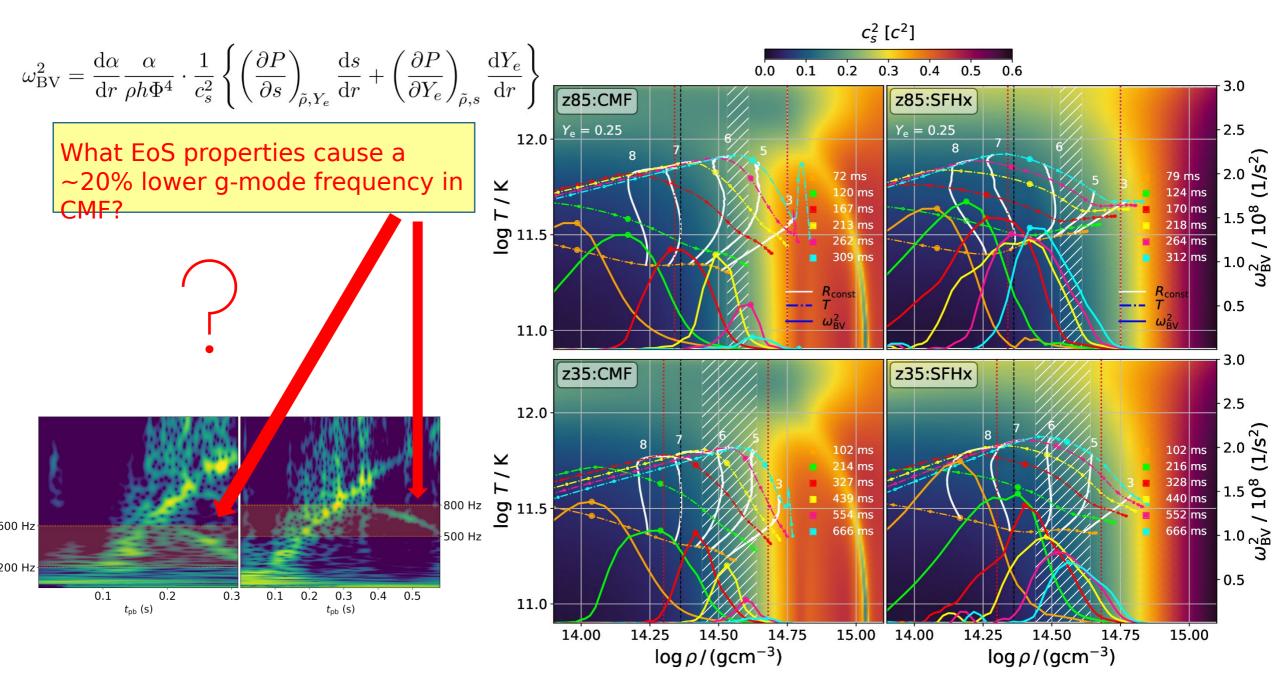
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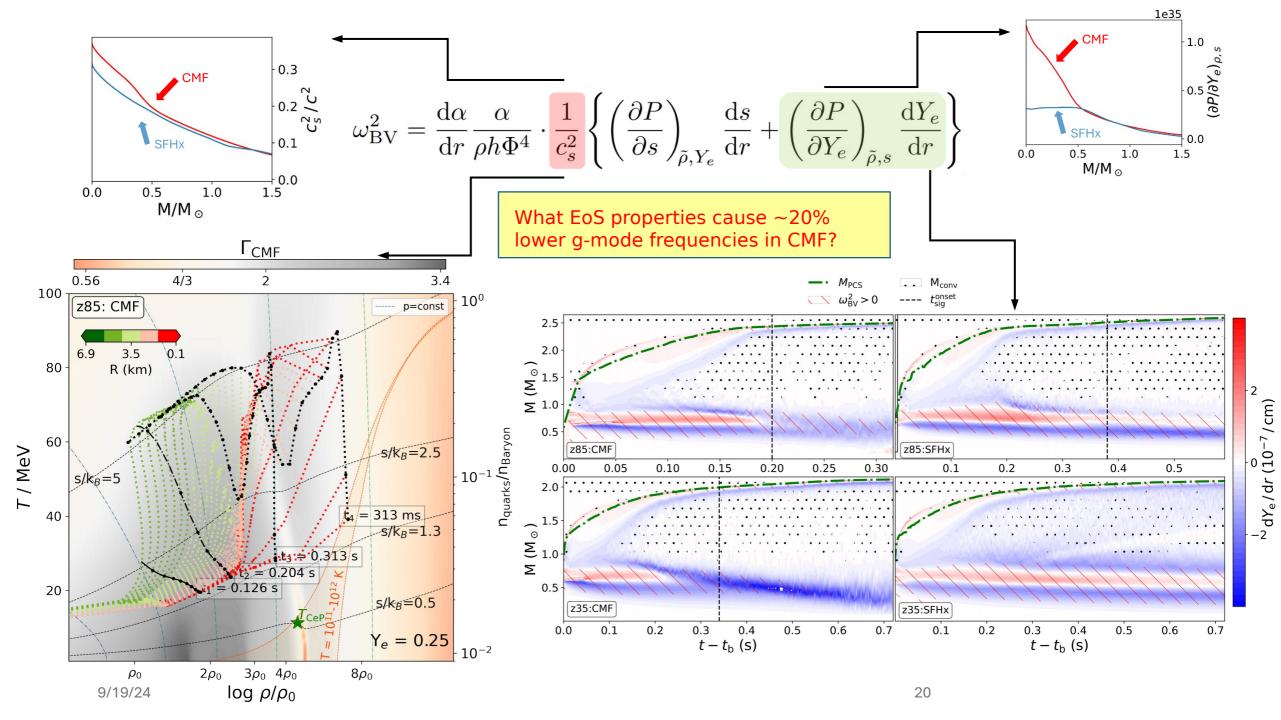


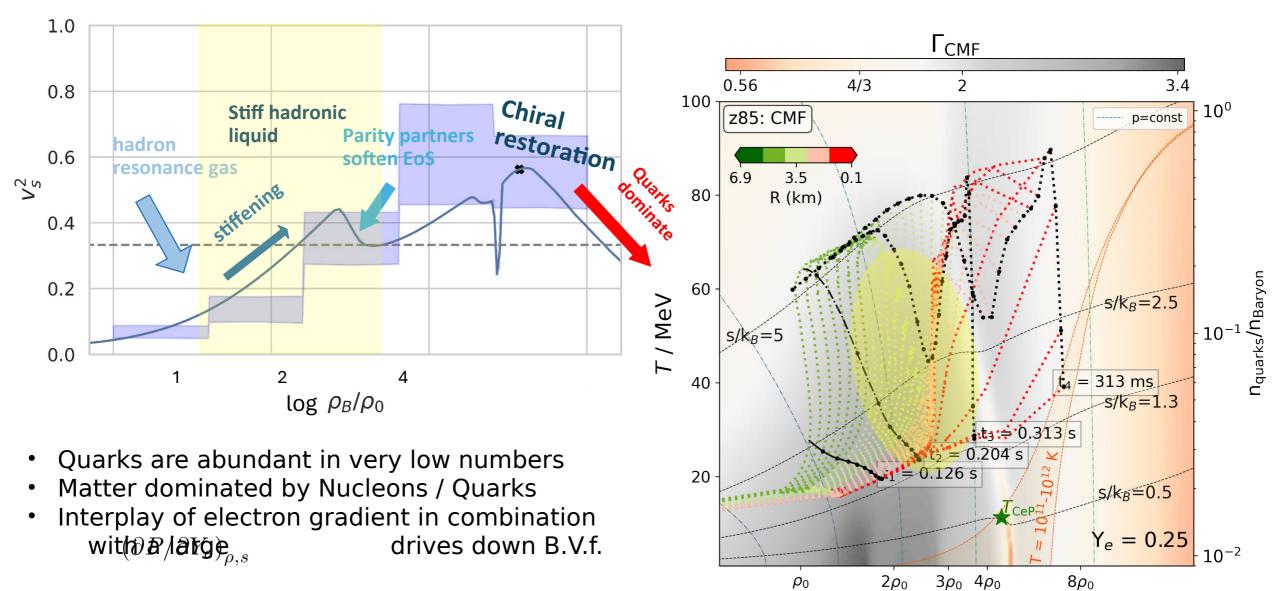
$$E_{\rm kin}^{\rm turb} = \frac{1}{2} \int_{\rm PCS} \delta v^2 \rho r^2 dV$$



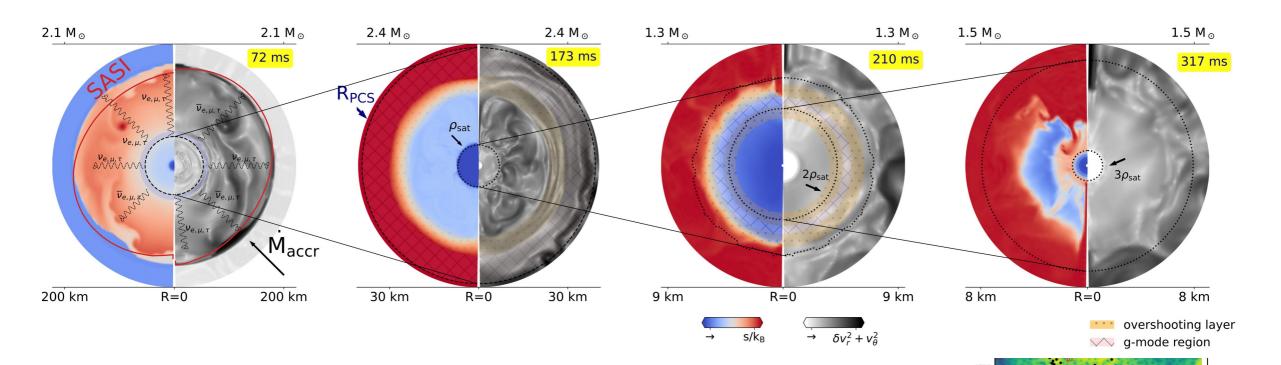








 $\log \rho/\rho_0$

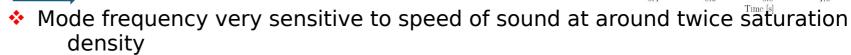


Summar y

Signal corresponds to g-mode oscillations in the PCS

~5-10 km (200-500 ms post-bounce)

I-mode [Zha et al., paper in prep]



Low frequency feature of mode lies within the sensitivity range of current GW detectors

2D versus 3D: Jakobus et al. 2024



of 1000

1bay 750

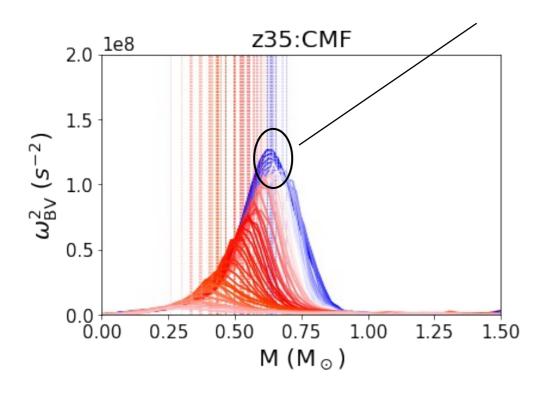
If a fluid parcel is perturbed- will it settle back into its equilibrium position?

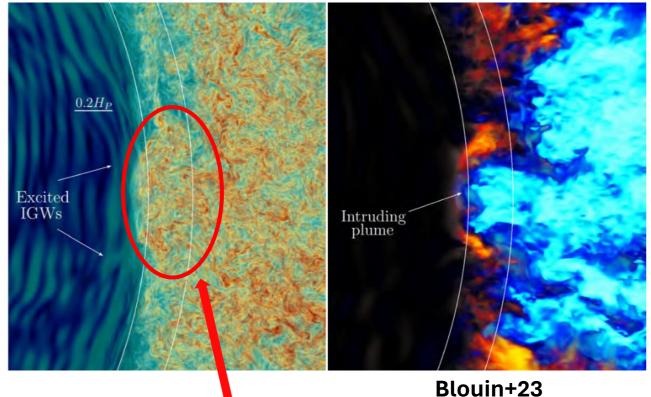
What also floats in water?

Oscillation of a displaced fluid element

0

Higher frequencies give rise to larger 2nd order time-derivatives of the mass quadrupole moment

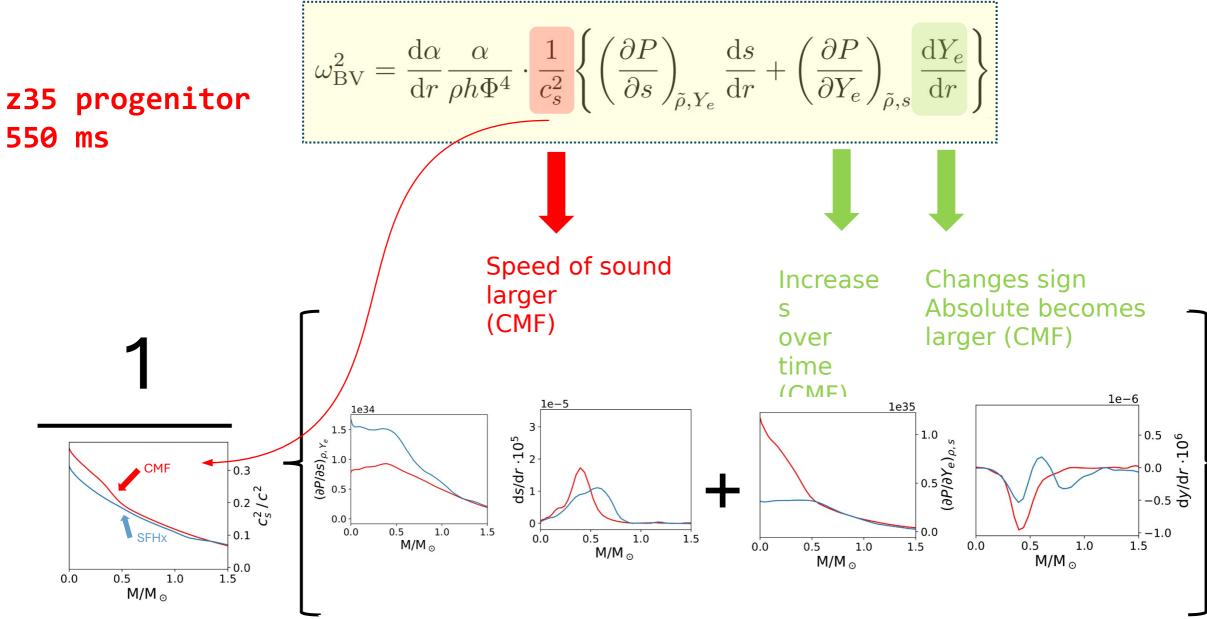




Based on the Ledoux criterion:

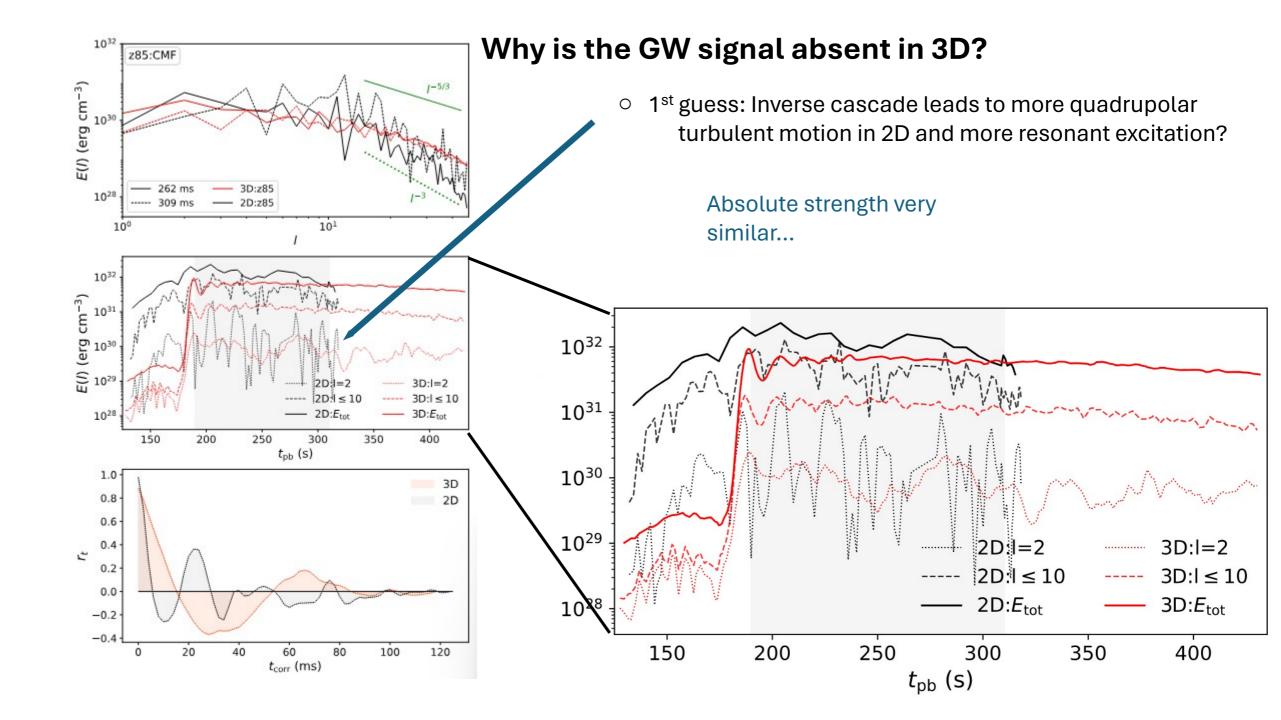
$$\omega_{\rm BV}^2 = \frac{\mathrm{d}\alpha}{\mathrm{d}r} \frac{\alpha}{\rho h \Phi^4} \cdot \frac{1}{c_s^2} \left\{ \left(\frac{\partial P}{\partial s} \right)_{\tilde{\rho}, Y_e} \frac{\mathrm{d}s}{\mathrm{d}r} + \left(\frac{\partial P}{\partial Y_e} \right)_{\tilde{\rho}, s} \frac{\mathrm{d}Y_e}{\mathrm{d}r} \right\}$$

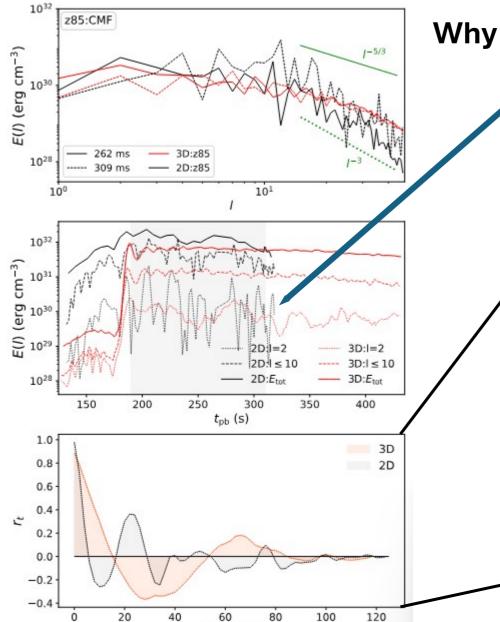
Overshoot into convectively stable region
Buoyancy force back toward convective region



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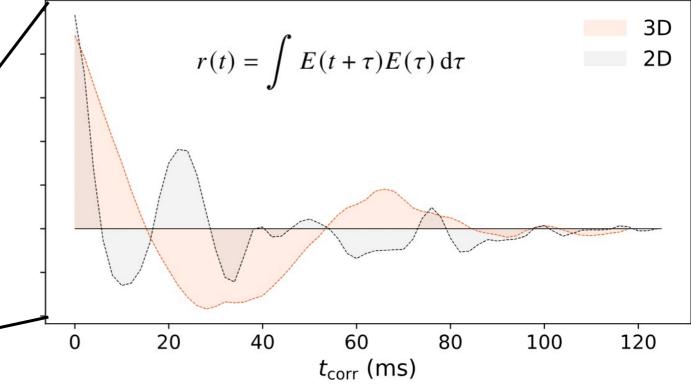
t_{corr} (ms)

Why is the GW signal absent in 3D?

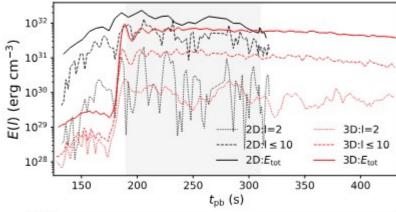
2nd guess: Larger temporal variability:

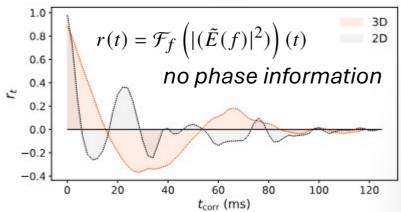
- Reflected in shorter auto-correlation time of E(2)
- Increased power at higher frequencies in the Fourier spectra

<u>Physical meaning:</u> faster decay rates of eddies in 2D greater velocity dispersion



10³² 285:CMF 10³⁰ 10²⁸ 10²⁸ 10²⁸ 10³⁰ 10³² 10³² 10³²





Why is the GW signal absent in 3D?

Autocorrelation function and Fourier transform look very different!

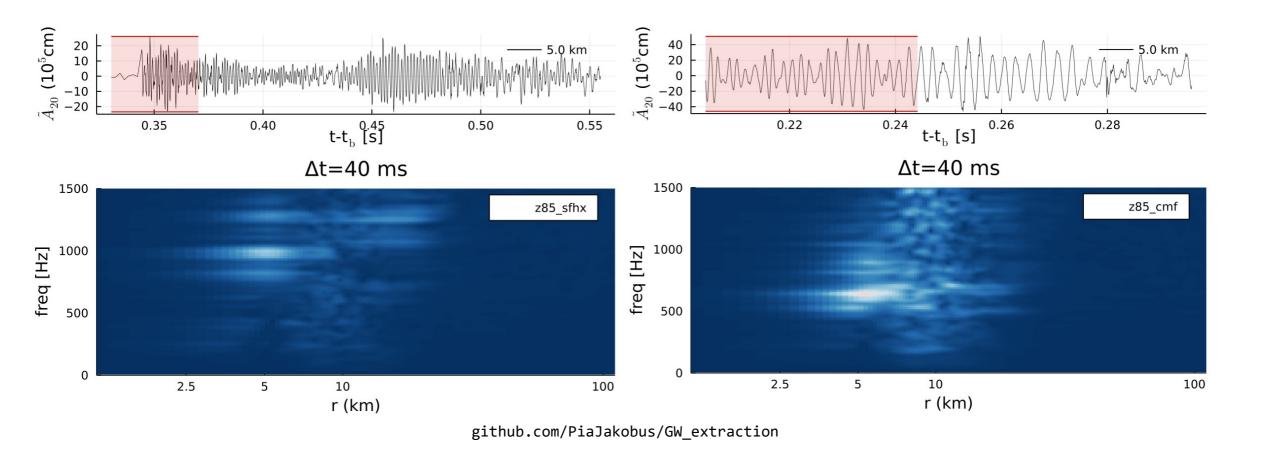
$$\tilde{E}(f) = 1/\sqrt{2\pi} \int E(t)e^{-2\pi i f t} dt \log \tilde{E}_f (\text{erg s cm}^{-3})$$

$$\frac{400}{285:3D}$$

$$\frac{300}{400}$$

$$\frac{285:2D}{400}$$

$$\frac{285:2D}{100}$$
See also Andresen+2017



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