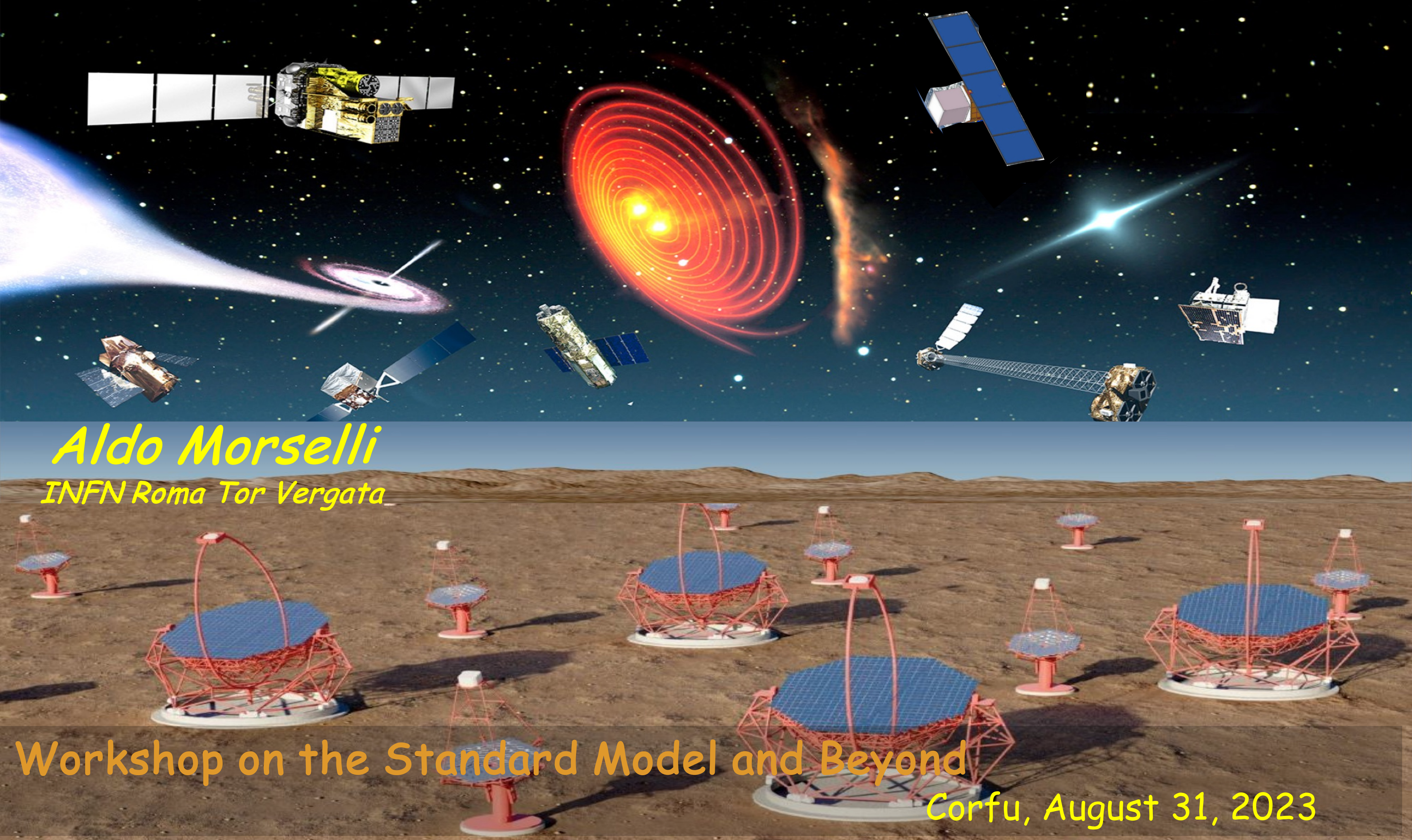
An update on Indirect dark-matter searches with gamma-rays experiments : status and future plans from 300 KeV to 100 TeV



Aldo Morselli
INFN Roma Tor Vergata

Workshop on the Standard Model and Beyond

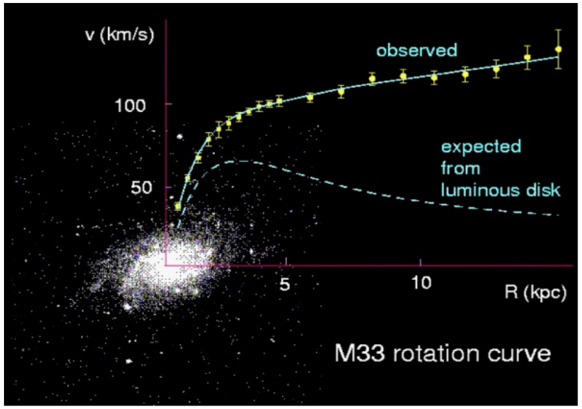
Corfu, August 31, 2023

Dark Matter EVIDENCE

In 1933, the astronomer Zwicky realized that the mass of the luminous matter in the Coma cluster was much smaller than its total mass implied by the motion of cluster member galaxies.

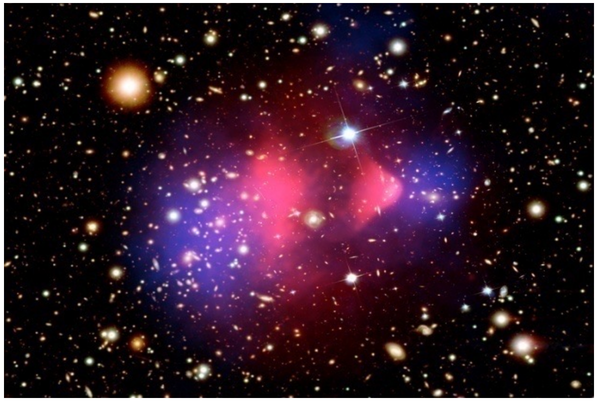
Since then, even more evidence:

Rotation curves of galaxies

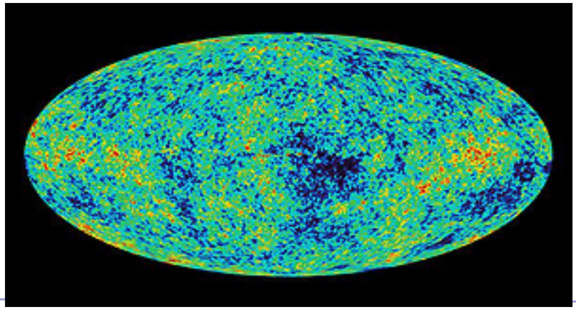


Gravitational lensing

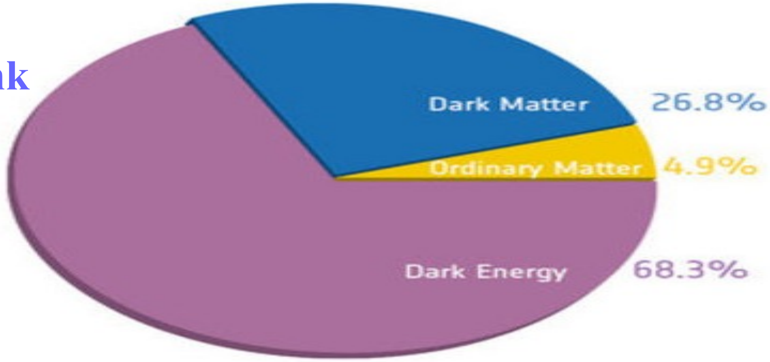
Bullet cluster



Structure formation as deduced from CMB



Data by Plank imply:



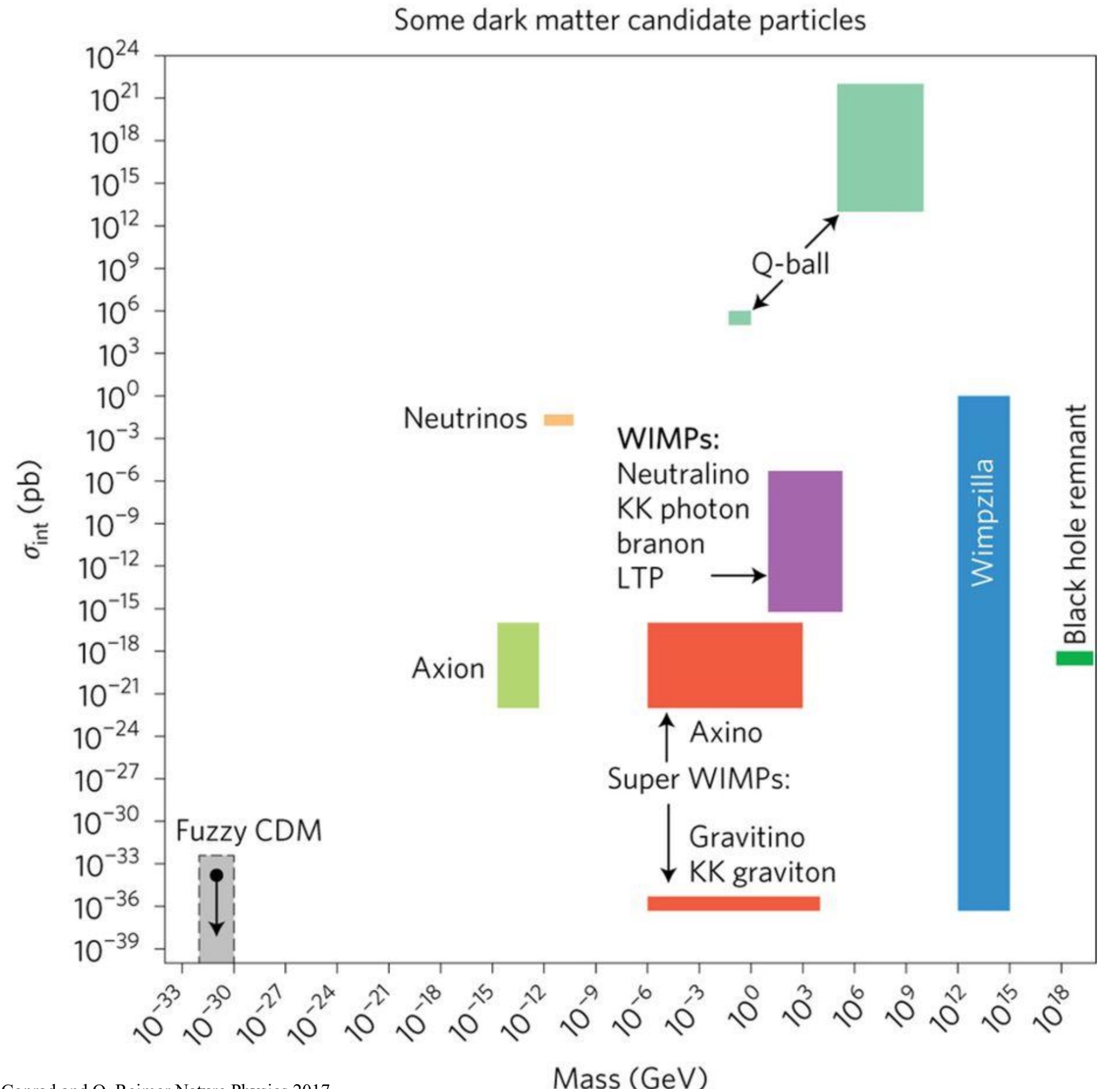
$\Omega_{DM} \approx 26.8\%$

$\Omega_{M} \approx 4.9\%$

Dark Matter

Dark Matter Candidates

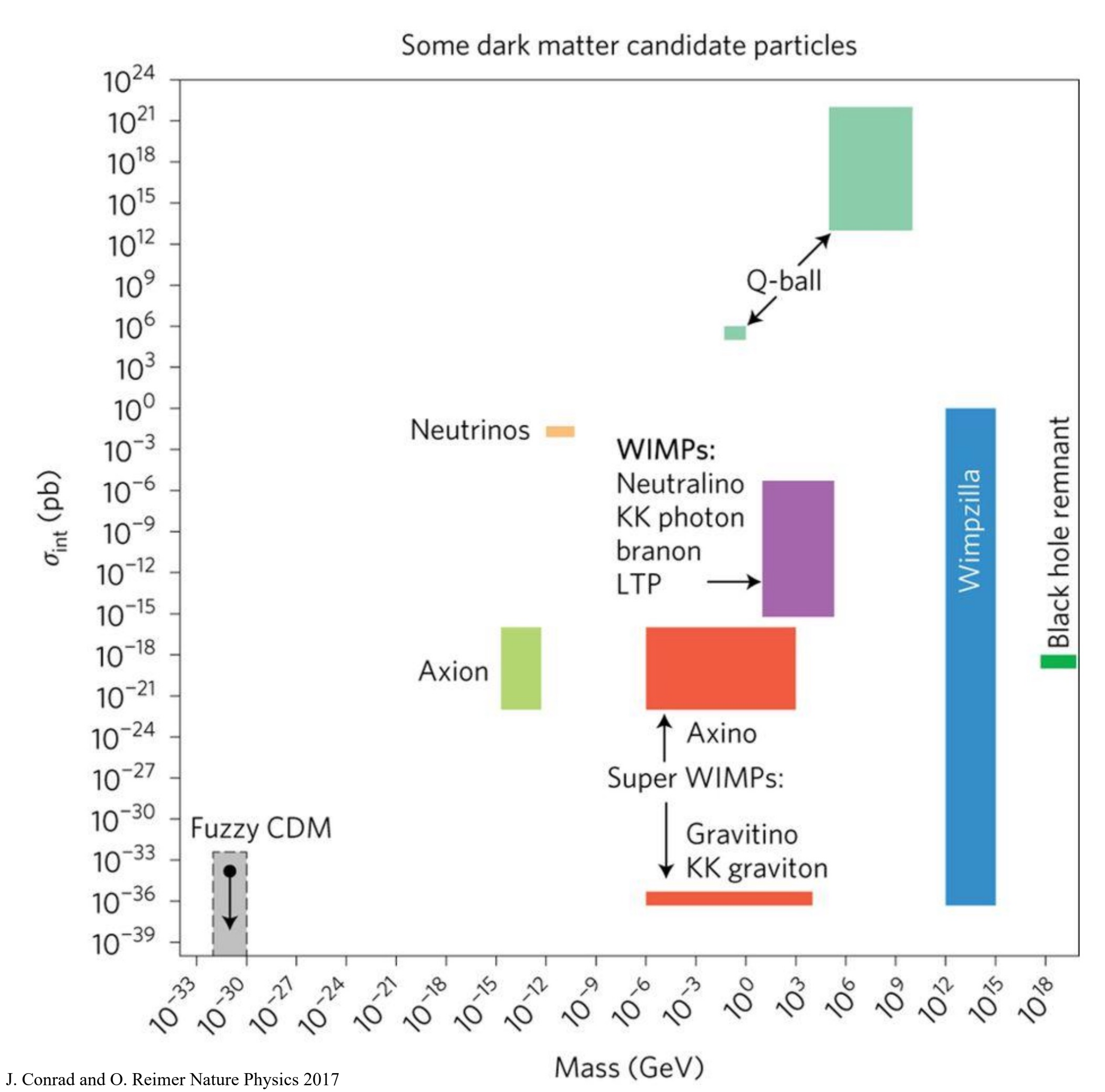
- Kaluza-Klein DM in UED
- Kaluza-Klein DM in RS
- Axion
- Axino
- Gravitino
- Photino
- SM Neutrino
- Sterile Neutrino
- Sneutrino
- Light DM
- Little Higgs DM
- Wimpzillas
- Q-balls
- Mirror Matter
- Champs (charged DM)
- D-matter
- Cryptons
- Self-interacting
- Superweakly interacting
- Braneworld DM
- Heavy neutrino
- WIMP
- Messenger States in GMSB
- Branons
- Chaplygin Gas
- Split SUSY
- Primordial Black Holes



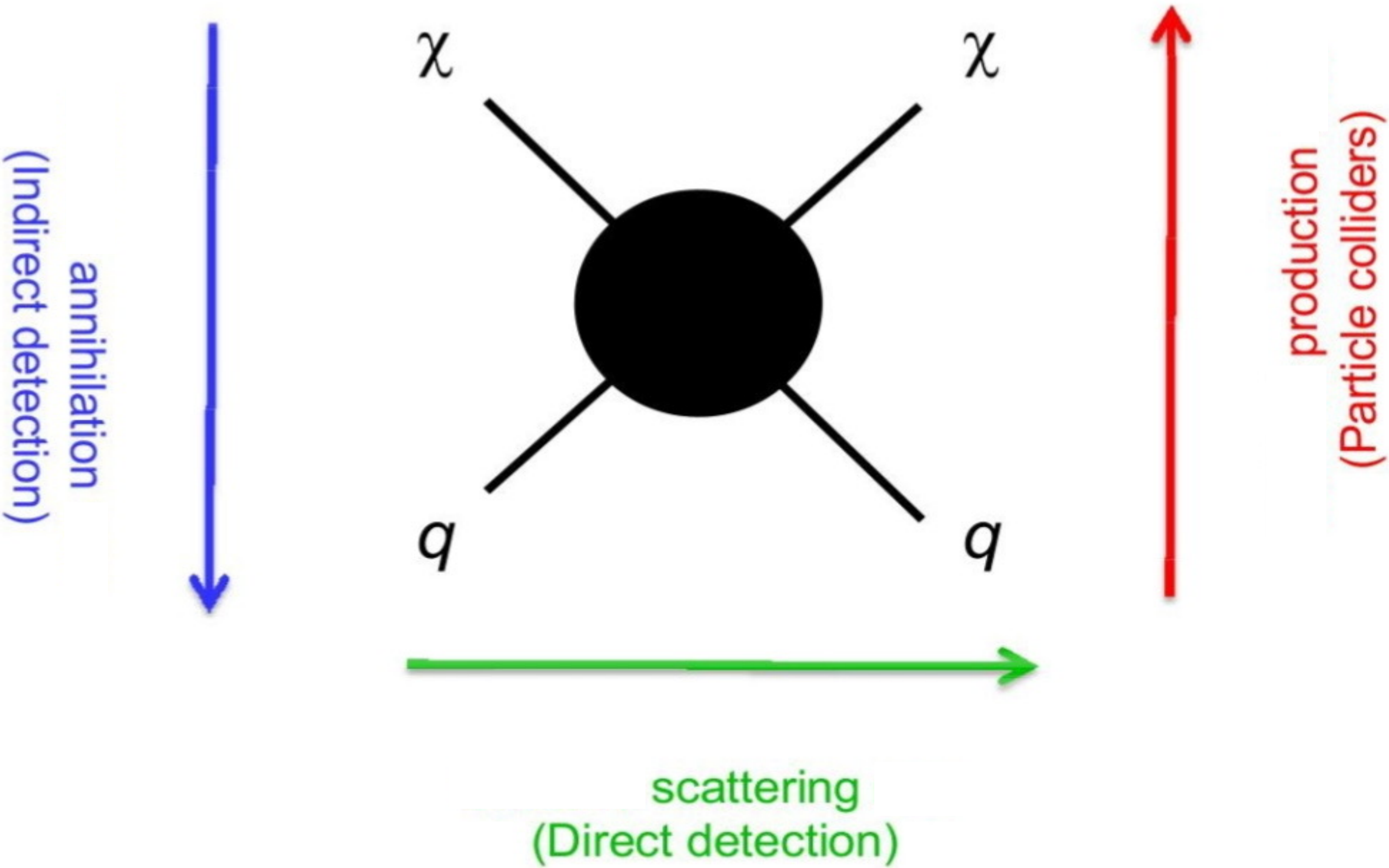
J. Conrad and O. Reimer Nature Physics 2017

Dark Matter Candidates

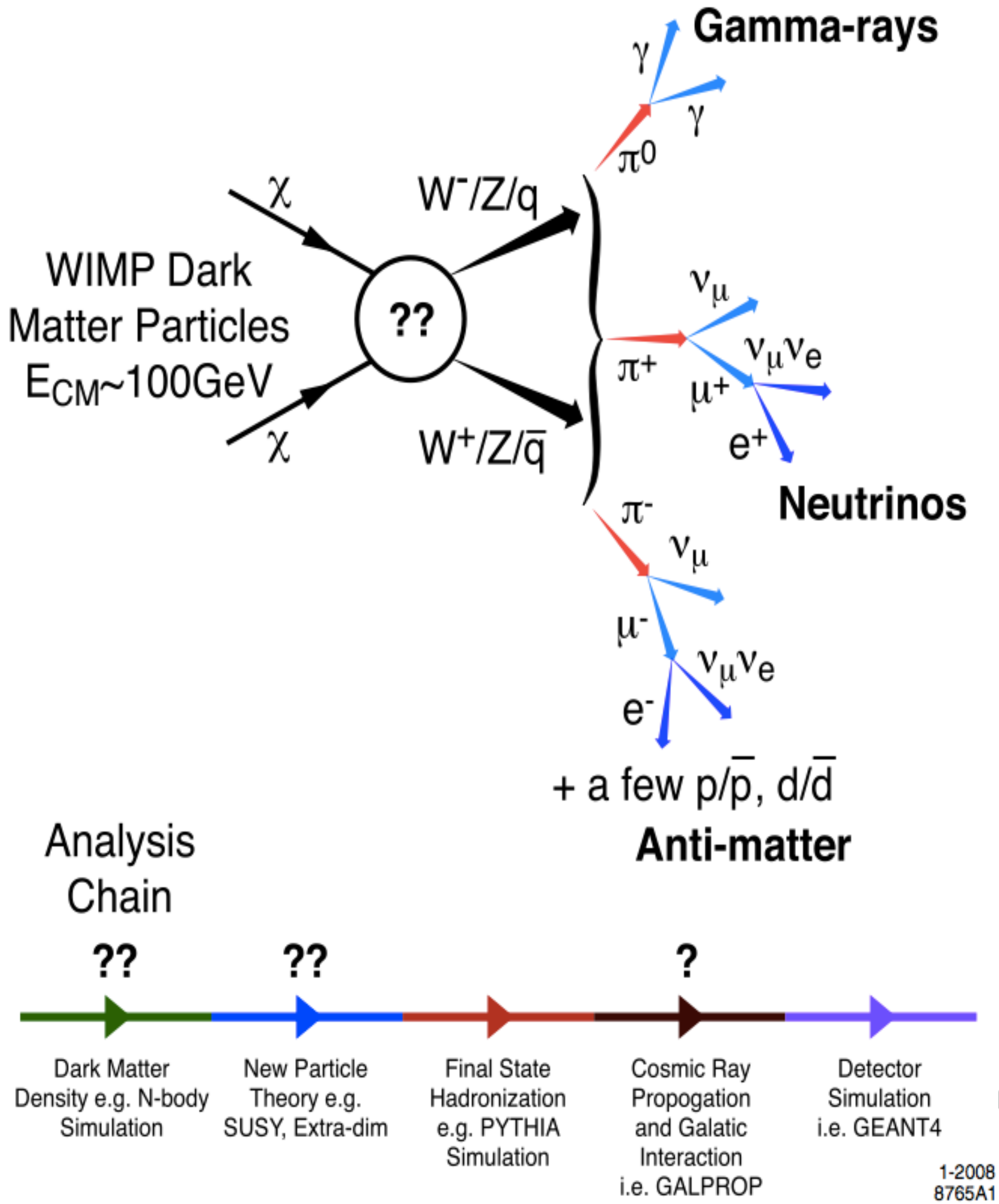
- Kaluza-Klein DM inUED
- Kaluza-Klein DM in RS
- Axion
- Axino
- Gravitino
- Photino
- SM Neutrino
- Sterile Neutrino
- Sneutrino
- Light DM
- Little Higgs DM
- Wimpzillas
- Q-balls
- Mirror Matter
- Champs (charged DM)
- D-matter
- Cryptons
- Self-interacting
- Superweakly interacting
- Braneworls DM
- Heavy neutrino
- **WIMP**
- Messenger States in GMSB
- Branons
- Chaplygin Gas
- Split SUSY
- Primordial Black Holes

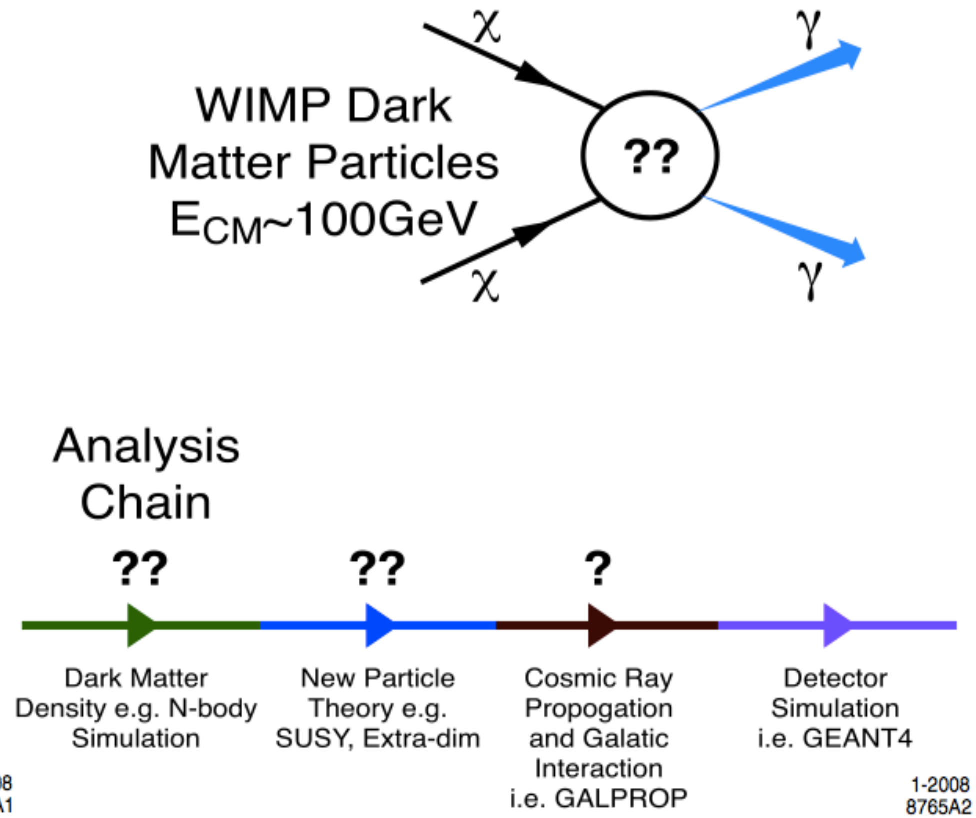


J. Conrad and O. Reimer Nature Physics 2017


annihilation
(Indirect detection)
χ
χ
q
q
production
(Particle colliders)
scattering
(Direct detection)

Annihilation channels



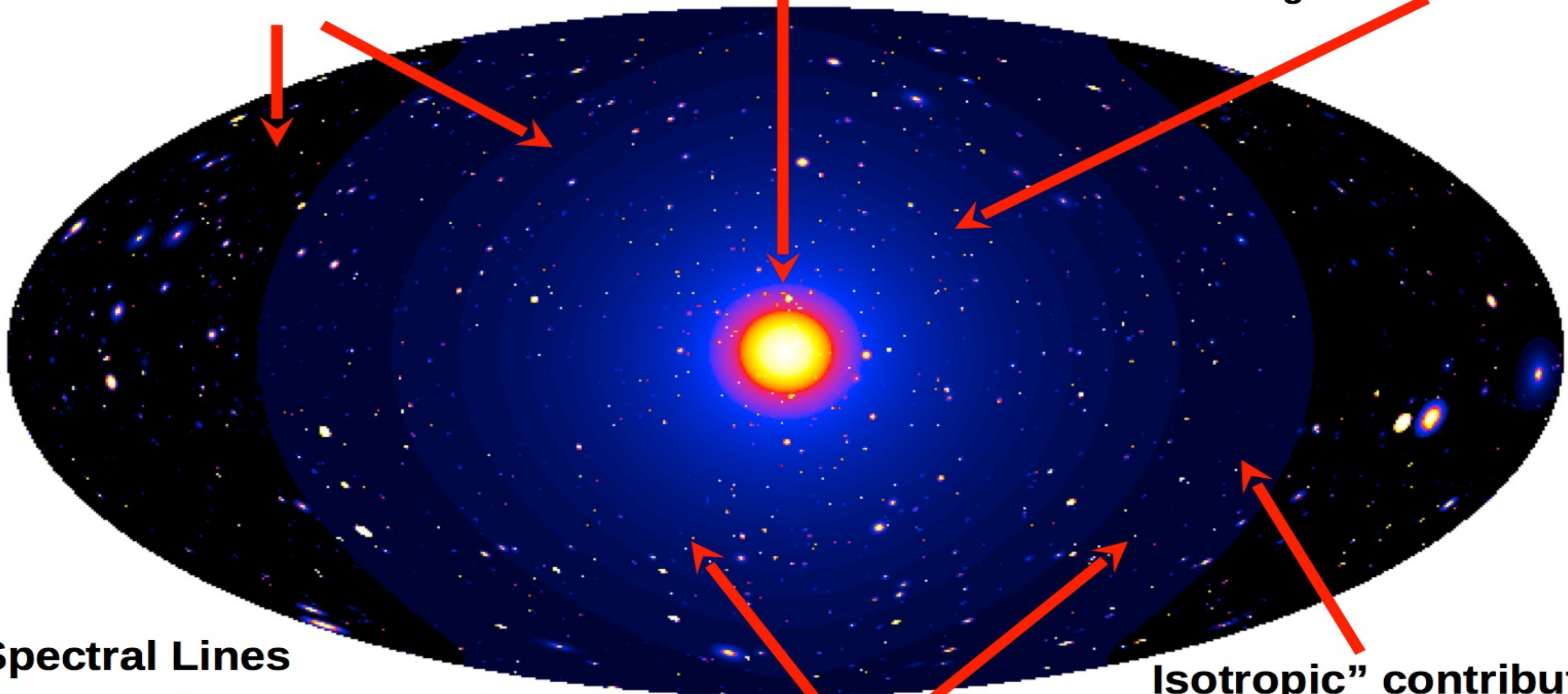


Dark Matter Search: Targets and Strategies

Satellites
Low background and good source id, but low statistics

Galactic Center
Good Statistics, but source confusion/diffuse background

Milky Way Halo
Large statistics, but diffuse background



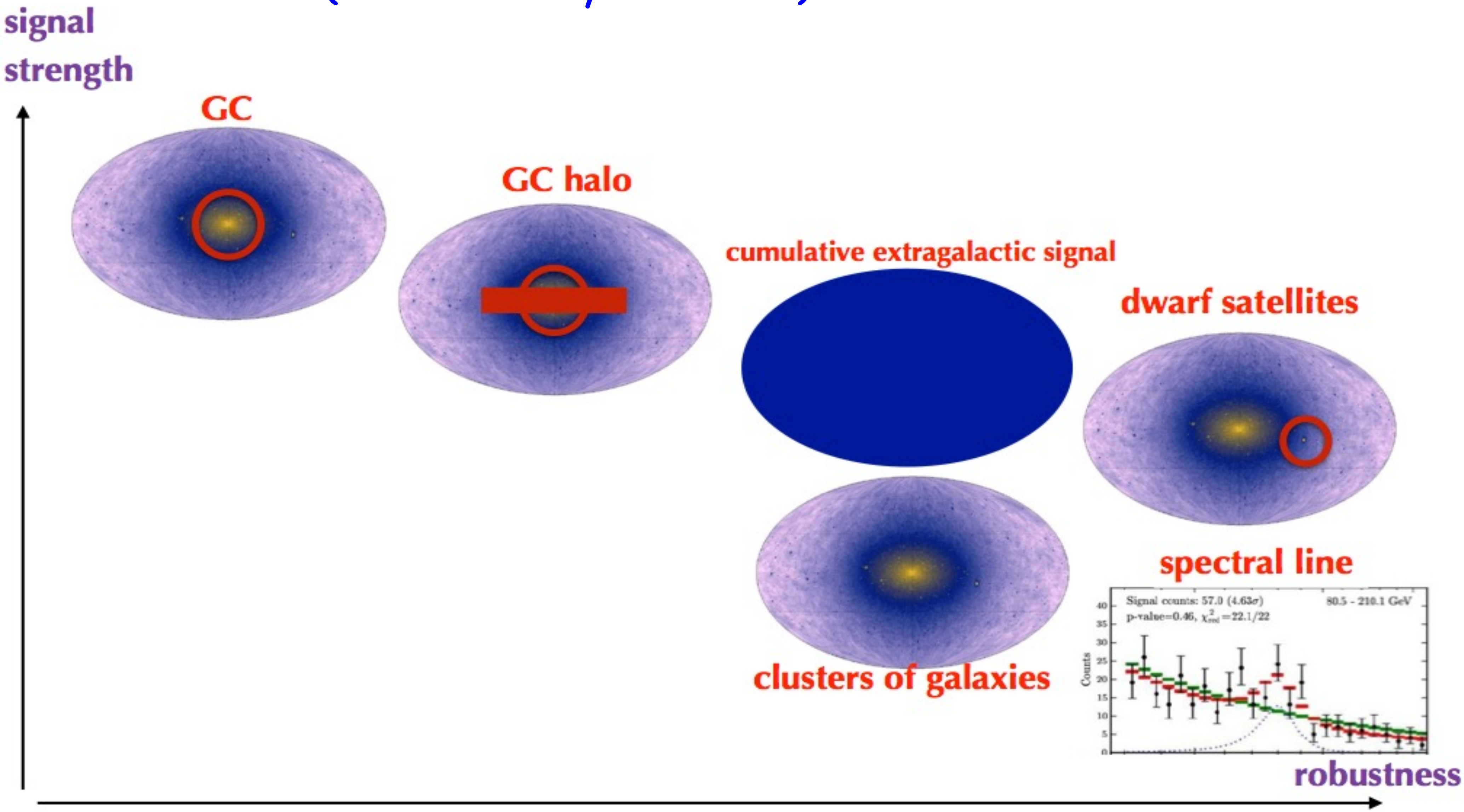
Spectral Lines
Little or no astrophysical uncertainties, good source id, but low sensitivity because of expected small branching ratio

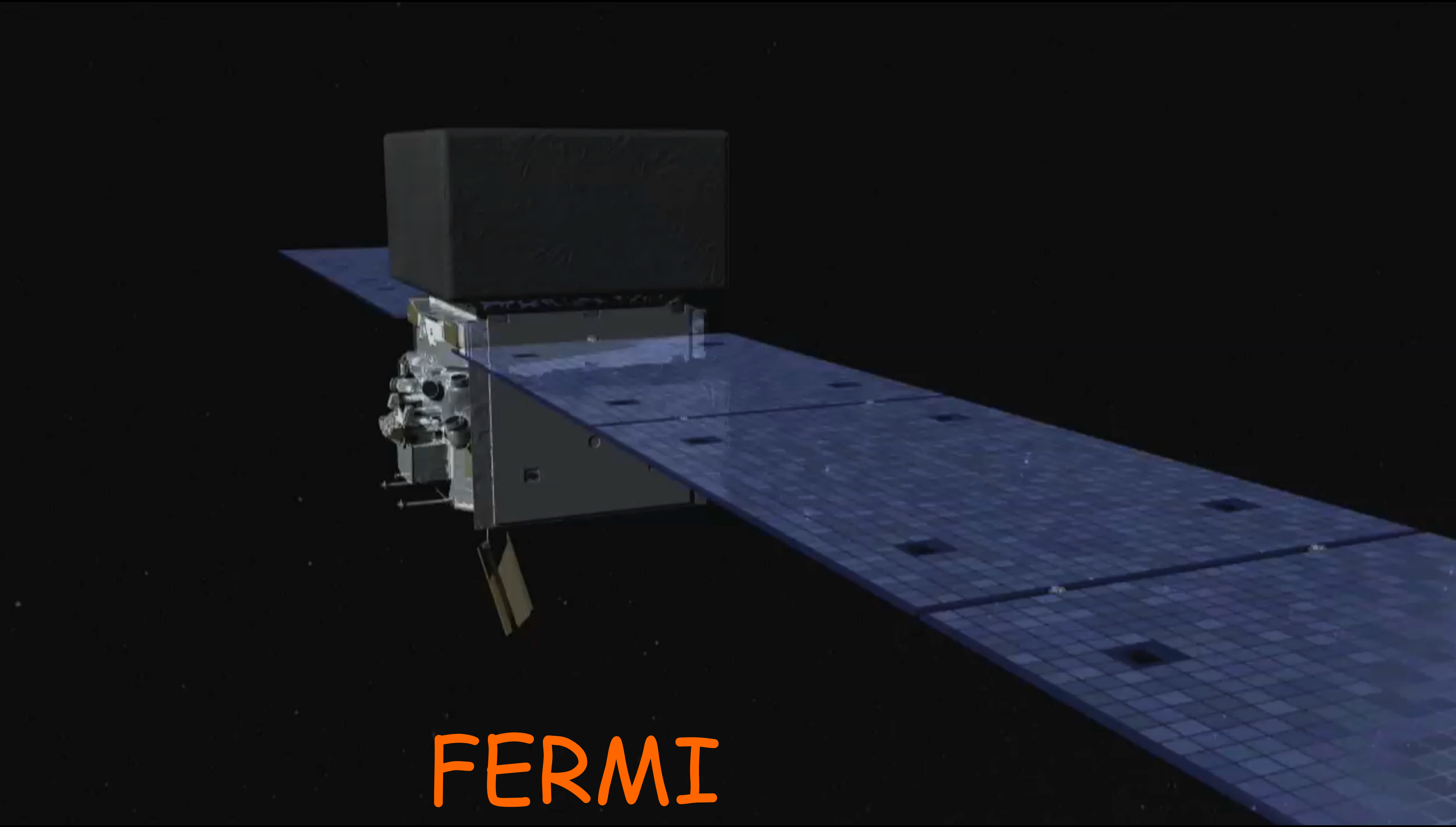
Isotropic" contributions
Large statistics, but astrophysics, galactic diffuse background

Galaxy Clusters
Low background, but low statistics

Dark Matter Search: Targets and Strategies

(Another way to see it)

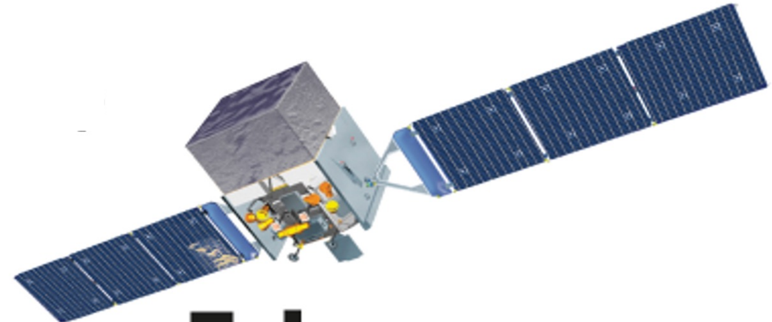


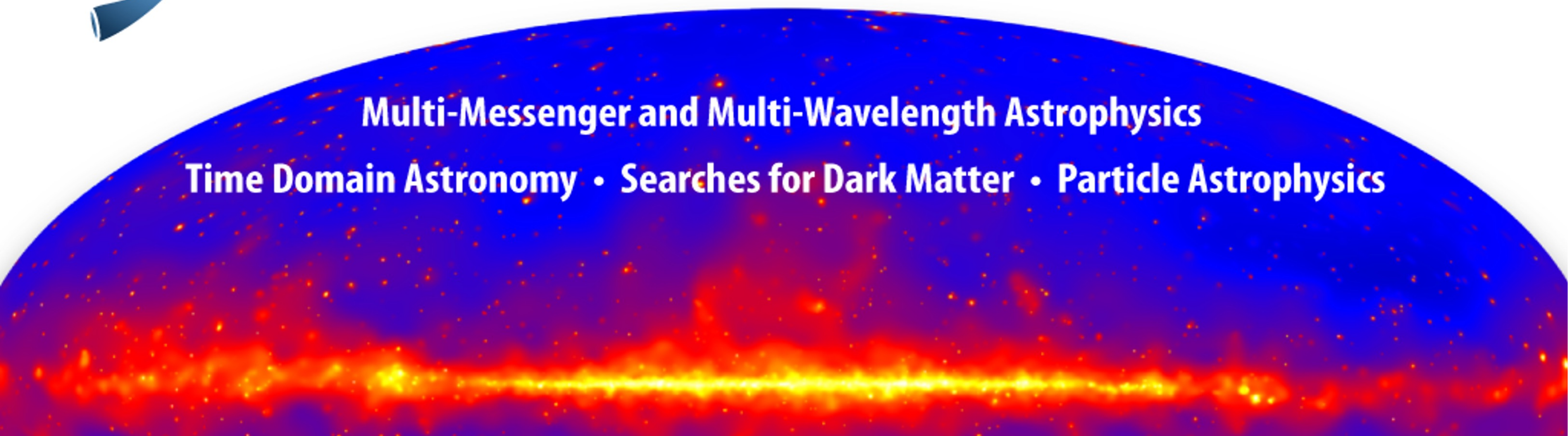


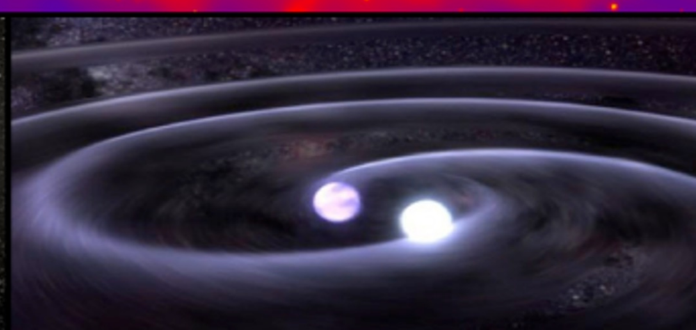
FERMI

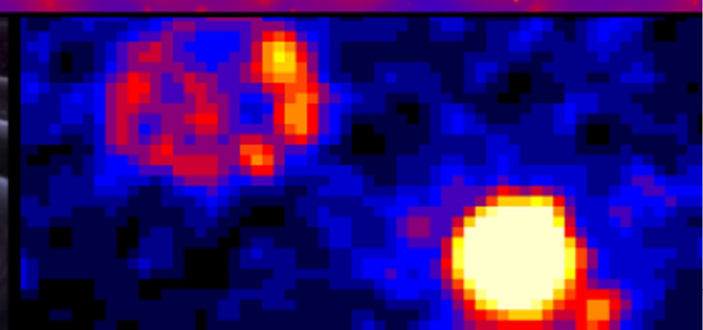
Large Area Telescope

Fermi Gamma-Ray Space Telescope



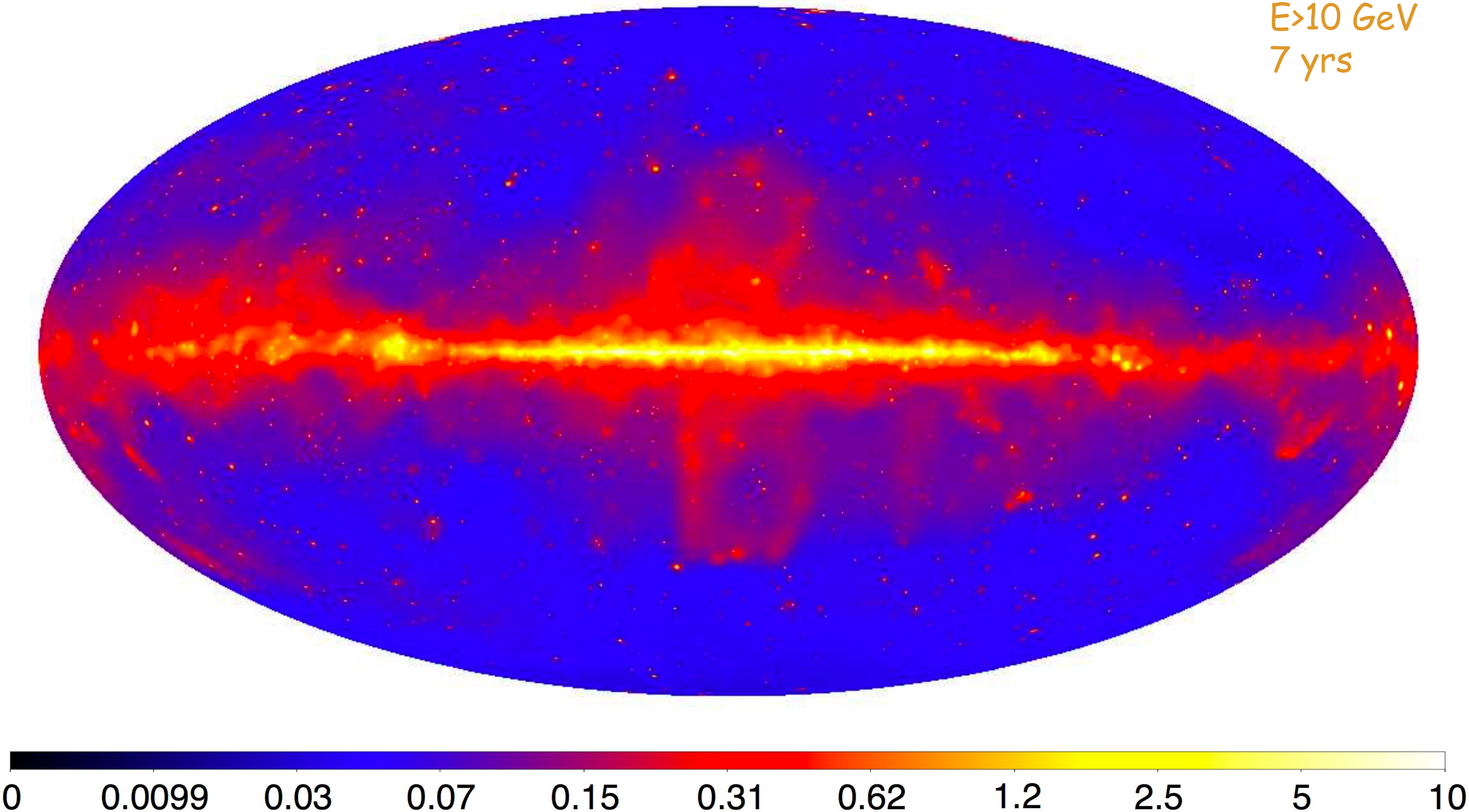






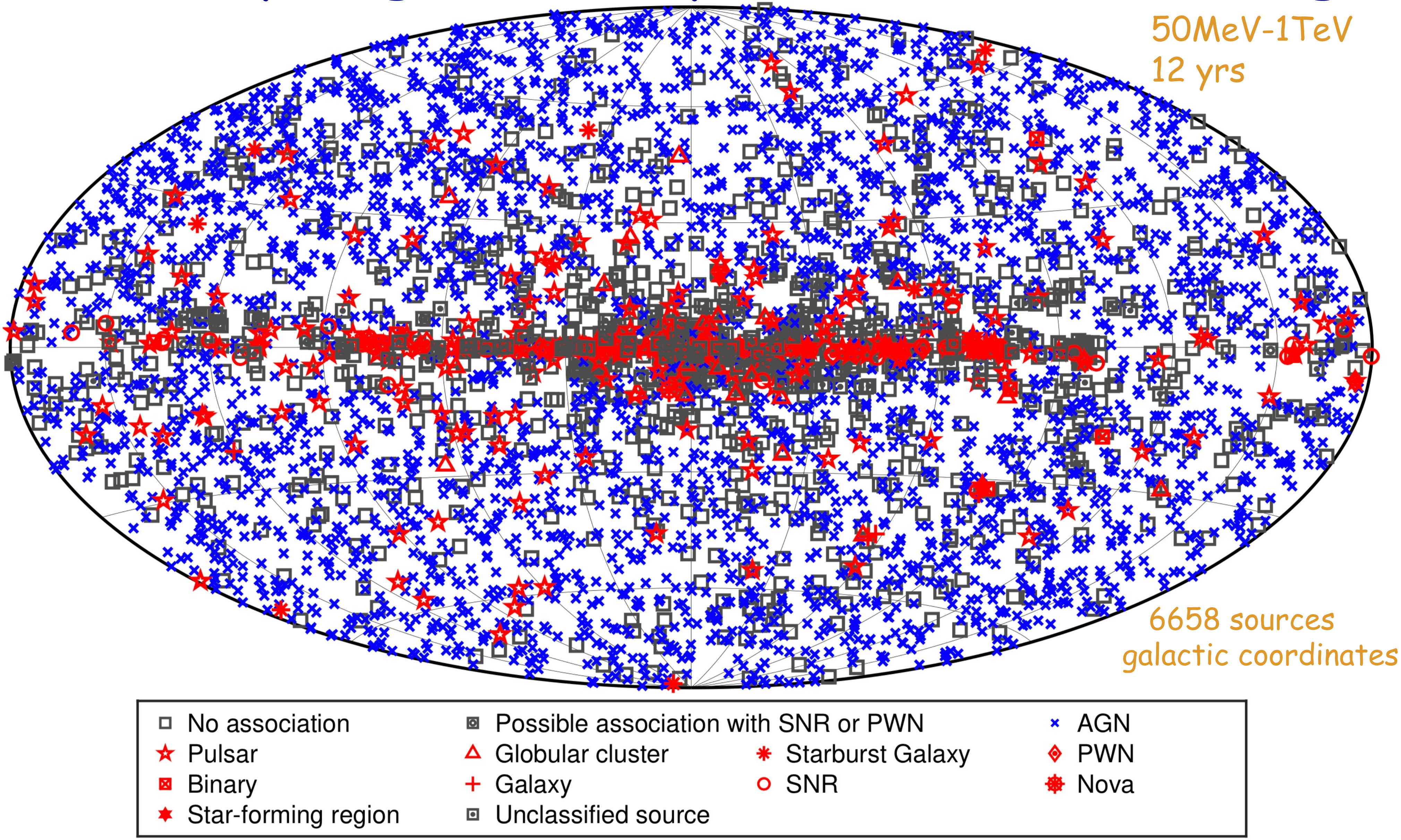
The sky in gamma-rays

E>10 GeV
7 yrs



M.Ackermann et al. [Fermi Coll.] 3FHL: The Third Catalog of Hard Fermi-LAT Sources ApJS 2017 232 arXiv:1702.00664

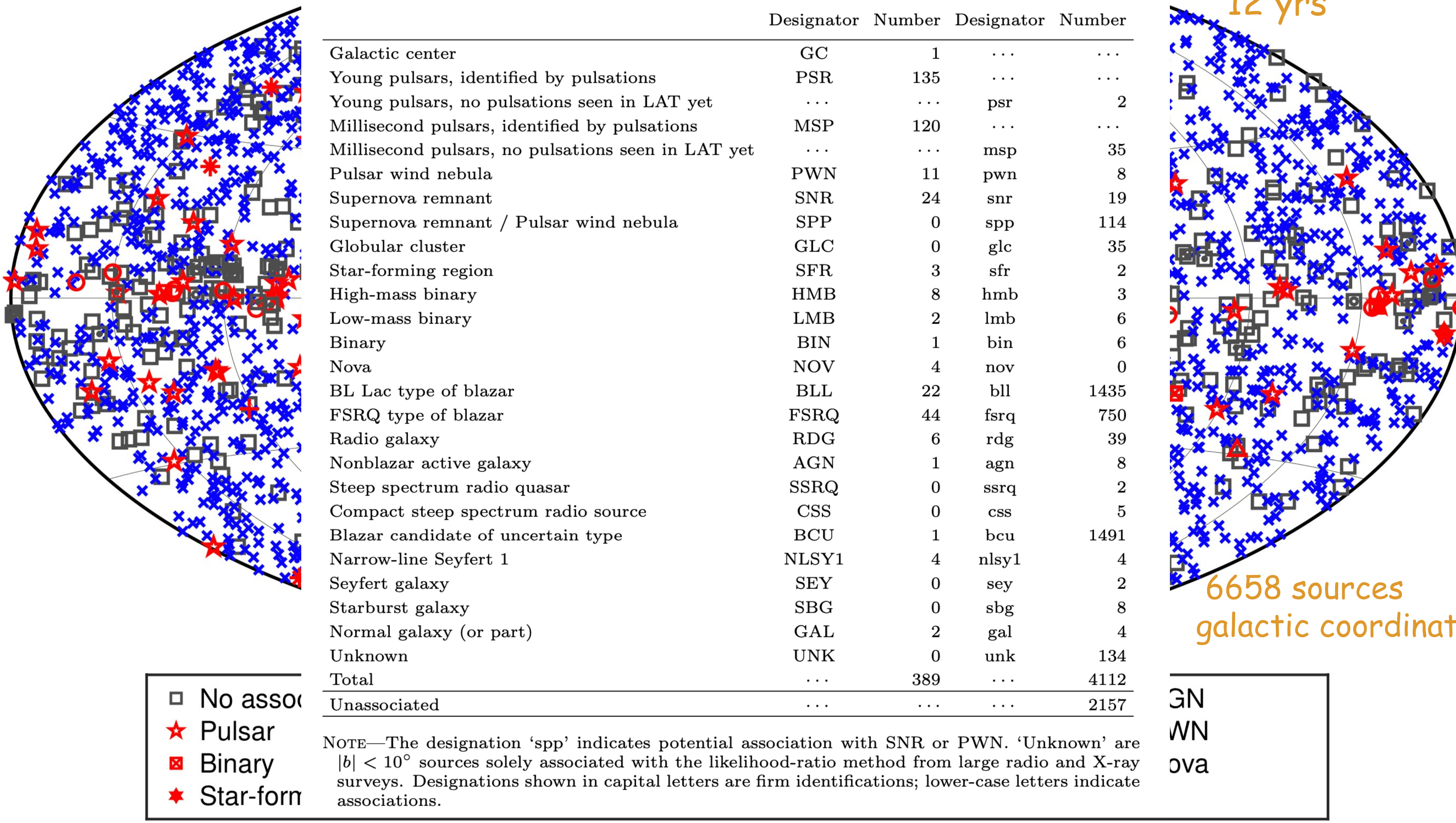
The sky in gamma-rays 4th source catalog



Incremental Fermi Fouth Source Catalog, ApJS 260, 53 (2022) arXiv: 2201.11184

The sky in gamma-rays 4th source catalog

50MeV-1TeV
12 yrs



Description	Identified		Associated	
	Designator	Number	Designator	Number
Galactic center	GC	1	...	...
Young pulsars, identified by pulsations	PSR	135	...	...
Young pulsars, no pulsations seen in LAT yet	...	...	psr	2
Millisecond pulsars, identified by pulsations	MSP	120	...	...
Millisecond pulsars, no pulsations seen in LAT yet	...	...	msp	35
Pulsar wind nebula	PWN	11	pwn	8
Supernova remnant	SNR	24	snr	19
Supernova remnant / Pulsar wind nebula	SPP	0	spp	114
Globular cluster	GLC	0	glc	35
Star-forming region	SFR	3	sfr	2
High-mass binary	HMB	8	hmb	3
Low-mass binary	LMB	2	lmb	6
Binary	BIN	1	bin	6
Nova	NOV	4	nov	0
BL Lac type of blazar	BLL	22	bll	1435
FSRQ type of blazar	FSRQ	44	fsrq	750
Radio galaxy	RDG	6	rdg	39
Nonblazar active galaxy	AGN	1	agn	8
Steep spectrum radio quasar	SSRQ	0	ssrq	2
Compact steep spectrum radio source	CSS	0	css	5
Blazar candidate of uncertain type	BCU	1	bcu	1491
Narrow-line Seyfert 1	NLSY1	4	nlsy1	4
Seyfert galaxy	SEY	0	sey	2
Starburst galaxy	SBG	0	sbg	8
Normal galaxy (or part)	GAL	2	gal	4
Unknown	UNK	0	unk	134
Total	...	389	...	4112
Unassociated	...	...	...	2157

Note—The designation ‘spp’ indicates potential association with SNR or PWN. ‘Unknown’ are $|b| < 10^\circ$ sources solely associated with the likelihood-ratio method from large radio and X-ray surveys. Designations shown in capital letters are firm identifications; lower-case letters indicate associations.

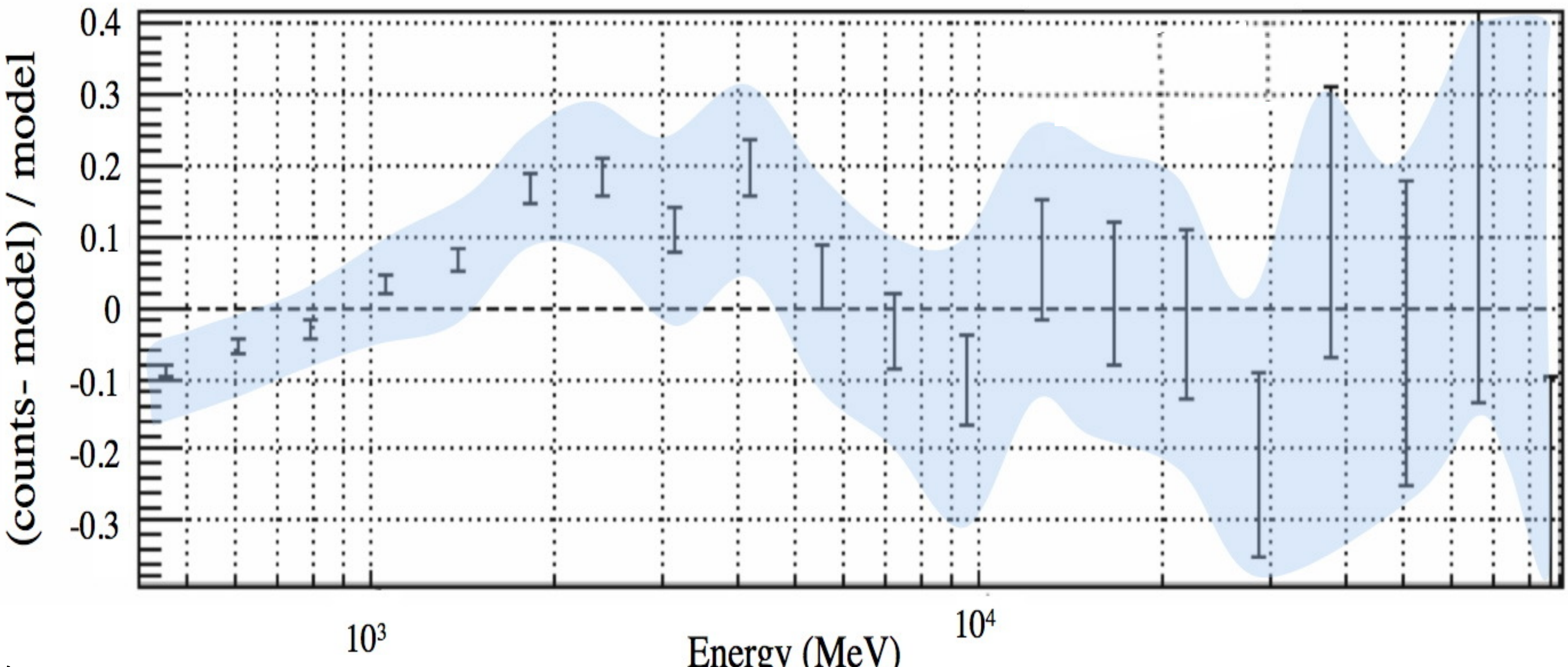
6658 sources
galactic coordinates

Incremental Fermi Fouth Source Catalog, ApJS 260, 53 (2022) arXiv: 2201.11184

The GeV excess 7° x7° region centered on the Galactic Center

11 months of data, E >400 MeV, front-converting events analyzed with binned likelihood analysis)

- The systematic uncertainty of the effective area (blue area) of the LAT is ~10% at 100 MeV, decreasing to 5% at 560 MeV and increasing to 20% at 10 GeV



V.Vitale, A.Morselli, Fermi Coll. 2009 arXiv:0912.3828 Fermi Symposium eConf Proceedings C091122

the GALACTIC CENTER : any hints of Dark Matter?

the beginning of the history :

The Galactic Center as a Dark Matter Gamma-Ray Source
A.Morselli, A. Lionetto, A. Cesarini, F. Fucito, P. Ullio, Nuclear Physics B 113B (2002) 213-220 [astro-ph/0211327]
A.Cesarini, F.Fucito, A.Lionetto, A.Morselli, P.Ullio Astroparticle Physics 21, 267-285, 2004 [astro-ph/0305075]

Possible Evidence For Dark Matter Annihilation In The Inner Milky Way From The Fermi Gamma Ray Space Telescope
Lisa Goodenough, Dan Hooper arXiv:0910.2998

Indirect Search for Dark Matter from the center of the Milky Way with the Fermi-Large Area Telescope
Vincenzo Vitale, Aldo Morselli, the Fermi/LAT Collaboration
Proceedings of the 2009 Fermi Symposium, 2-5 November 2009, eConf Proceedings C091122 **arXiv:0912.3828 21 Dec 2009**

Search for Dark Matter with Fermi Large Area Telescope: the Galactic Center
V.Vitale, A.Morselli, the Fermi-LAT Collaboration NIM A 630 (2011) 147–150 (Available online 23 June 2010)

Dark Matter Annihilation in The Galactic Center As Seen by the Fermi Gamma Ray Space Telescope
Dan Hooper, Lisa Goodenough. (21 March 2011). 21 pp. **Phys.Lett. B697 (2011) 412-428**
....

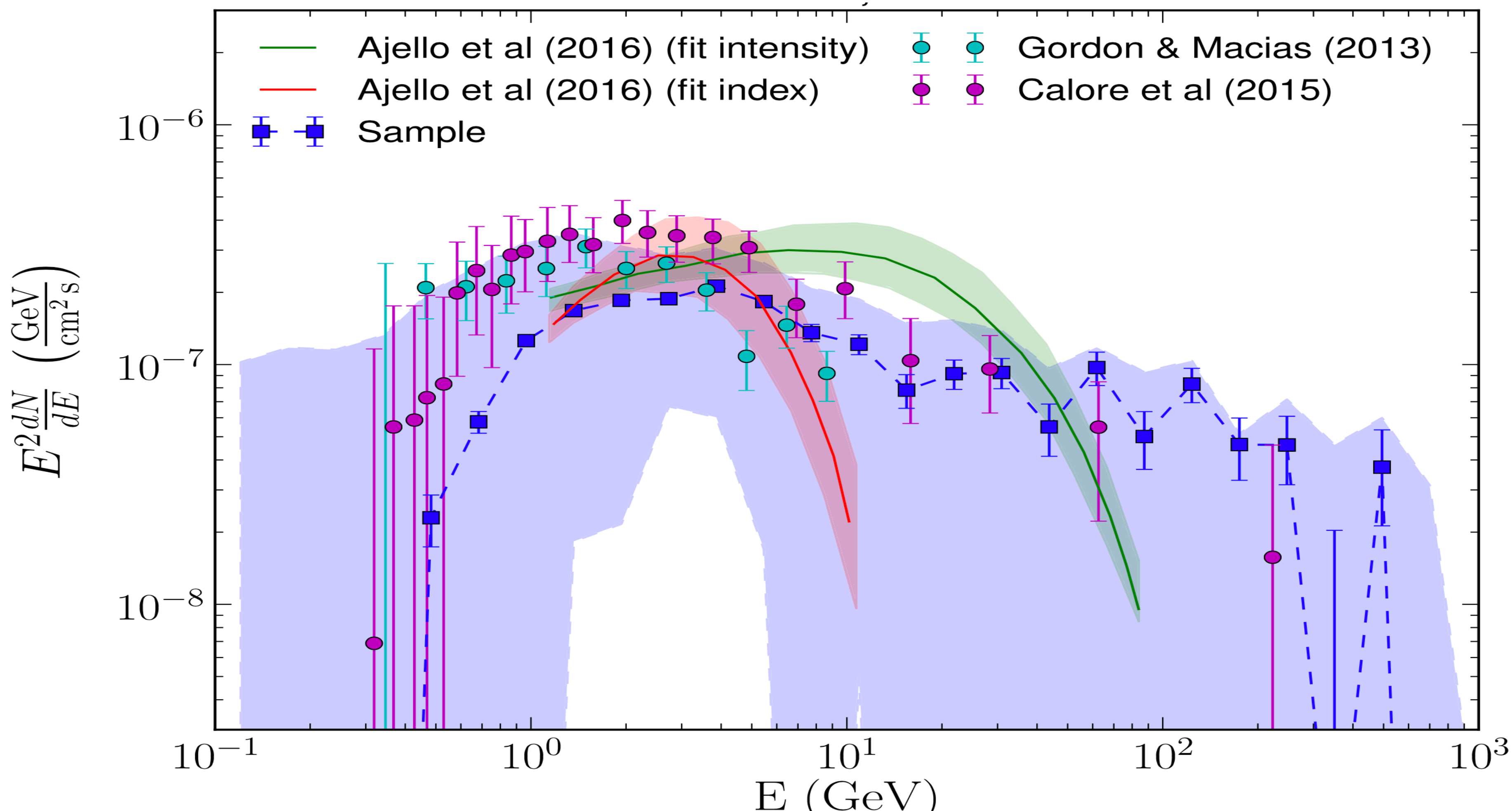
Background model systematics for the Fermi GeV excess
F.Calore, I. Cholis, C. Weniger JCAP03(2015)038 **arXiv:1409.0042v1**

Fermi–LAT observations of high-energy γ-ray emission toward the galactic centre
M. Ajello et al.[Fermi-LAT Coll.] **Apj 819:44 2016 arXiv:1511.02938**

The Fermi galactic center GeV excess and implications for dark matter
M. Ajello et al.[Fermi-LAT Coll.] **Apj 819:44 2016 arXiv:1511.02938**

Revisiting the Gamma-Ray Galactic Center Excess with Multi-Messenger Observations
IC, Zhong, McDermott, Surdutovich, PRD 105, 103023 (2022)

The GeV excess (Pass8 analysis)



following uncertainties have relatively small effect on the excess spectrum

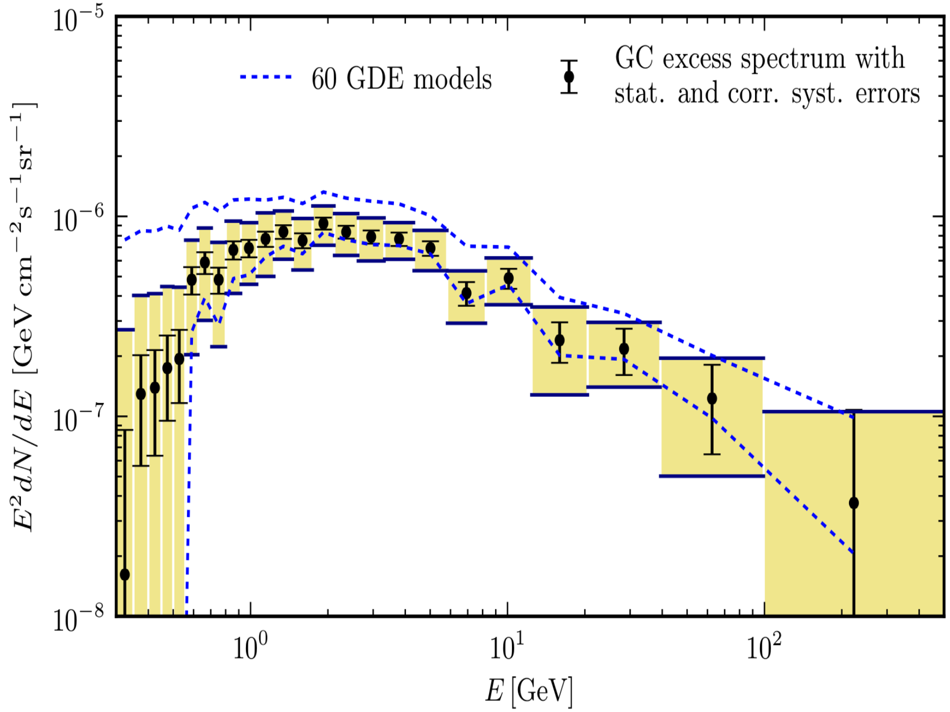
- Variation of GALPROP models – Distribution of gas along the line of sight

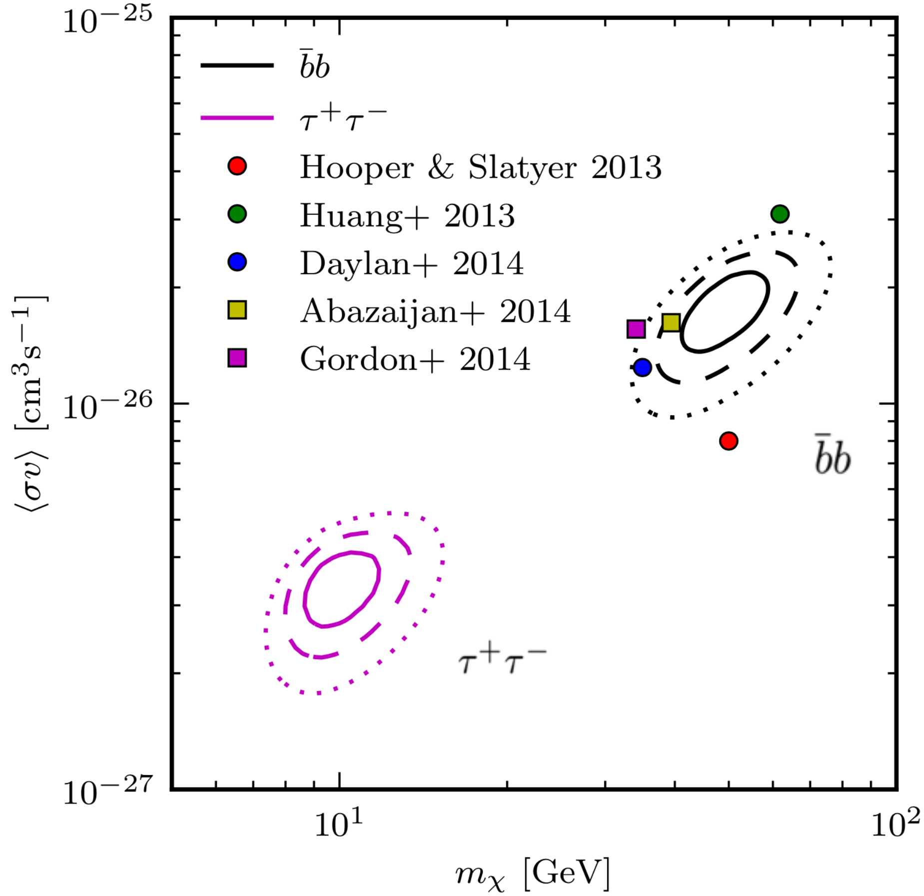
- **Most significant sources of uncertainty are:**

- Fermi bubbles morphology at low latitude – Sources of CR electrons near the GC

Fermi-LAT Collaboration Apj 840:43 2017 May 1 arXiv:1704.03910

The GeV excess





A lot of activity outside the Fermi collaboration with claims of evidence for dark matter in the Galactic Center

Calore et al., arXiv:1409.0042
Cholis et al., Phys. Rev. D 105, 103023 (2022) arXiv:2112.09706

The GeV excess : Other explanations exist

- past activity of the Galactic center

(e.g. Petrovic et al., arXiv:1405.7928, Carlson & Profumo arXiv:1405.7685)

- Series of Leptonic Cosmic-Ray Outbursts

Cholis et al. arXiv:1506.05119

- Stellar population of the X-bulge and the nuclear bulge

Macias et al. arXiv:1611.06644

- Population of pulsars in the Galactic bulge

e.g. , Yuan and Zhang arXiv:1404.2318v1, Lee et al. arXiv:1506.05124, Bartels et.al. 1506.05104

M.Ajello et al. [Fermi-LAT Coll.] Phys. Rev. D 95, 082007 (2017) [arXiv:1704.07195]

....

How to discriminate between different hypothesis ?

How to discriminate between different hypothesis ?

eROSITA

- Modeling of the Fermi bubbles
- Look for correlated features near the Galactic center

HESS, MAGIC, CTA

- Fermi bubbles near the GC are much brighter
- Possible to see with Cherenkov telescopes?

Radio observations, MeerKAT, SKA

- Search for individual pulsars in the halo around the GC

Radio surveys, Planck

- Look for correlated synchrotron emission near the GC

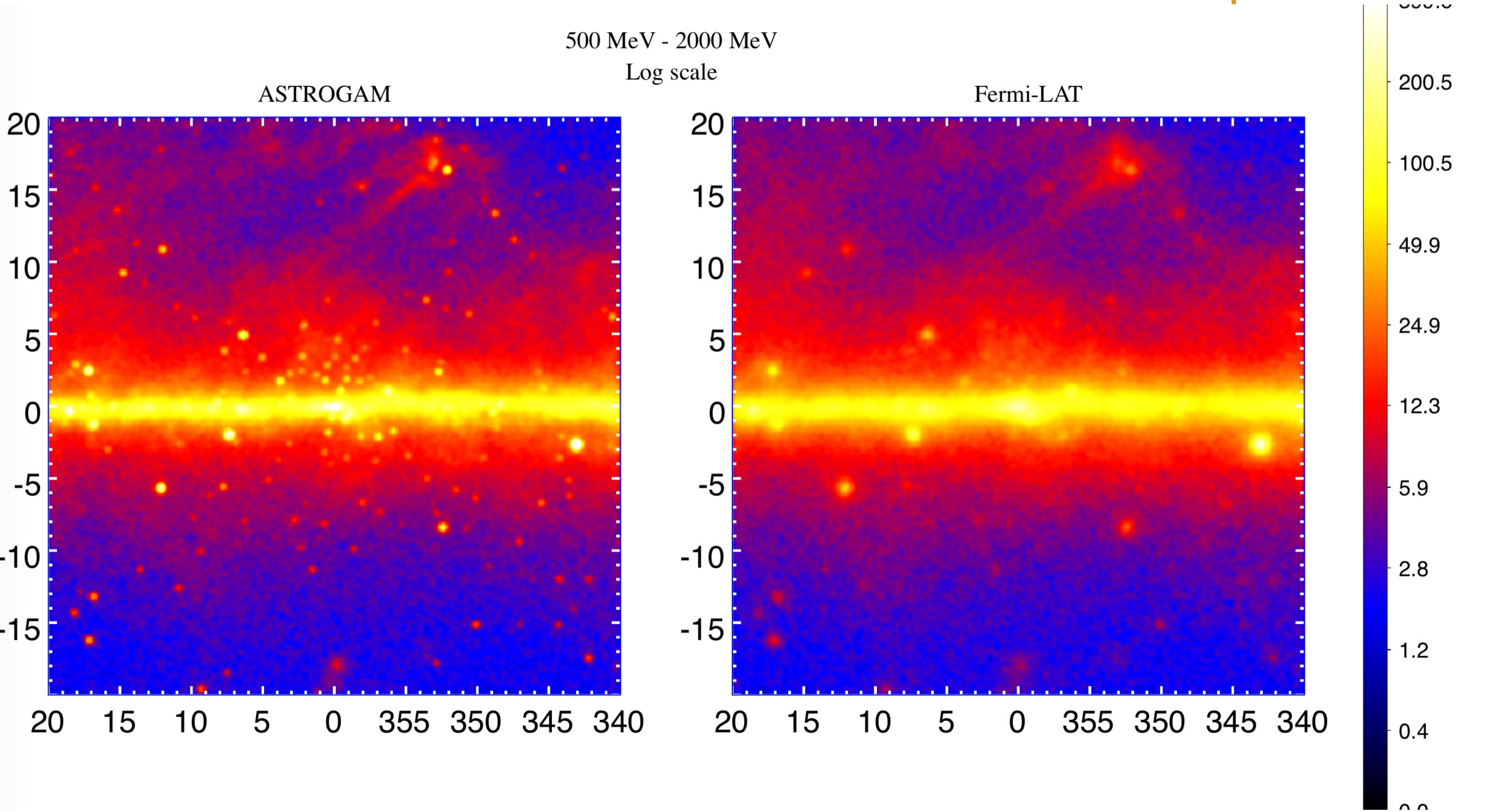
More Fermi LAT analysis

- Diffuse emission modeling
- Analysis of point sources near the GC

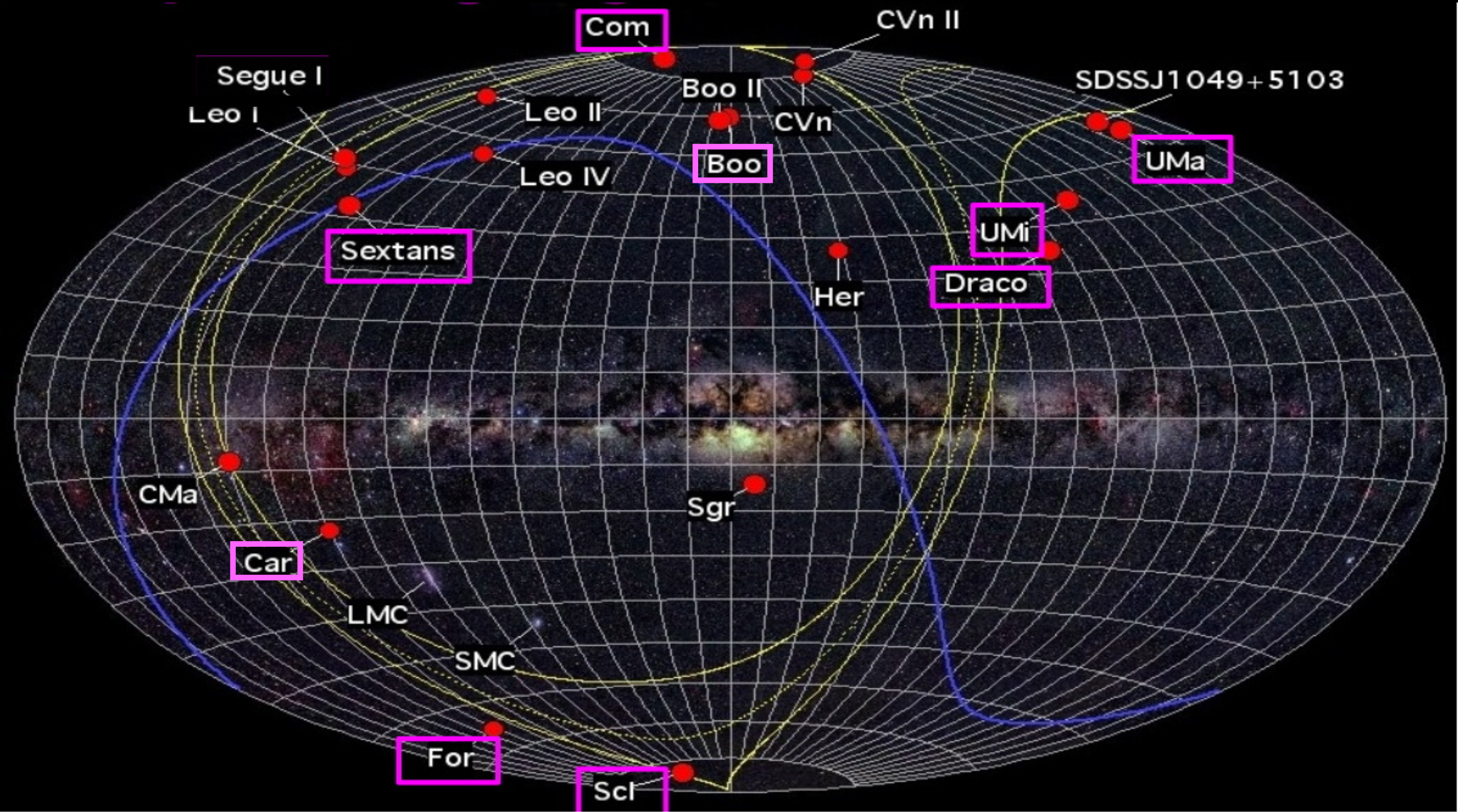
But ultimately We need a new experiment with better angular resolution below 100 MeV

Galactic Center Region 0.5-2 GeV

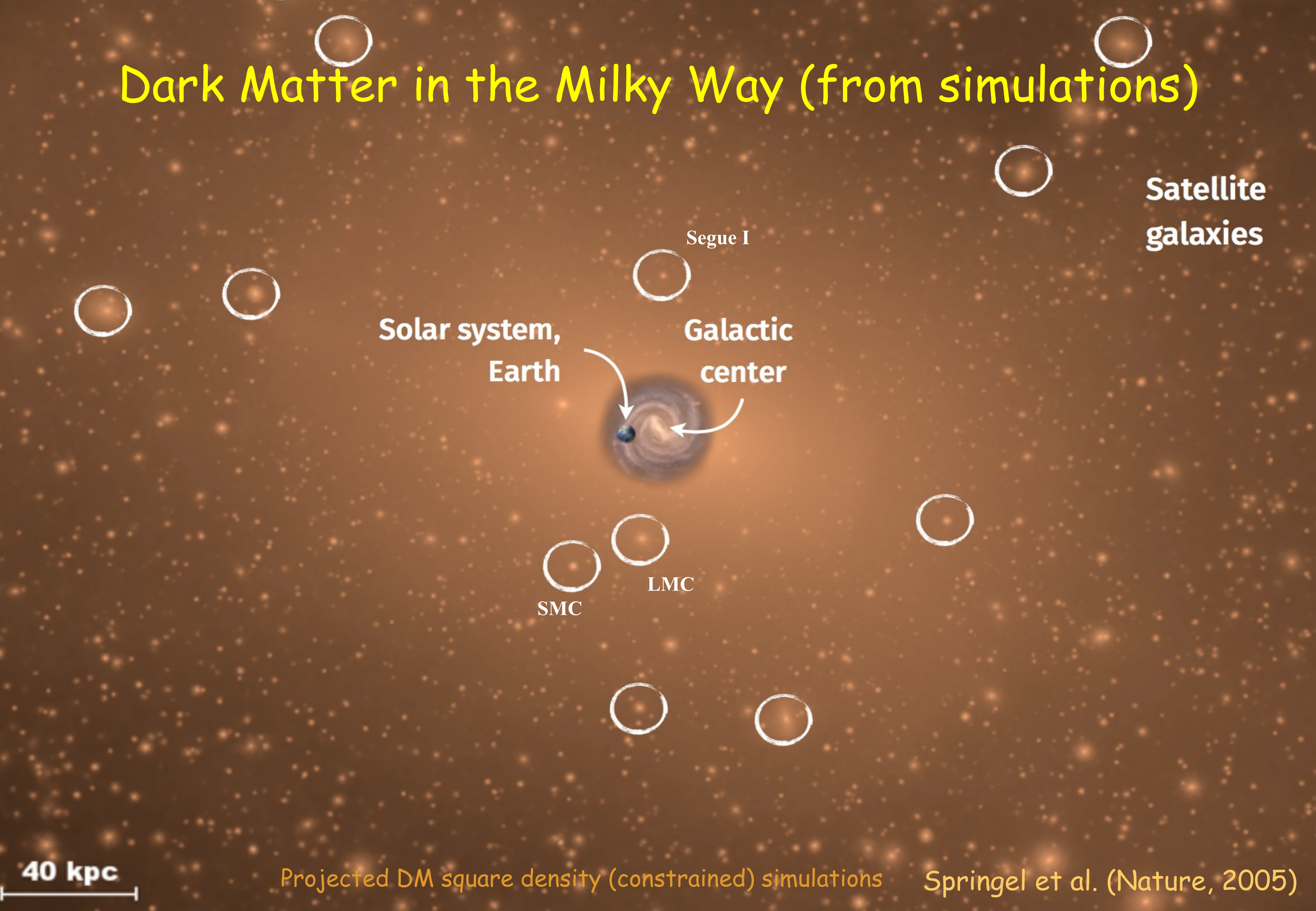
Fermi PSF Pass7 rep v15 source



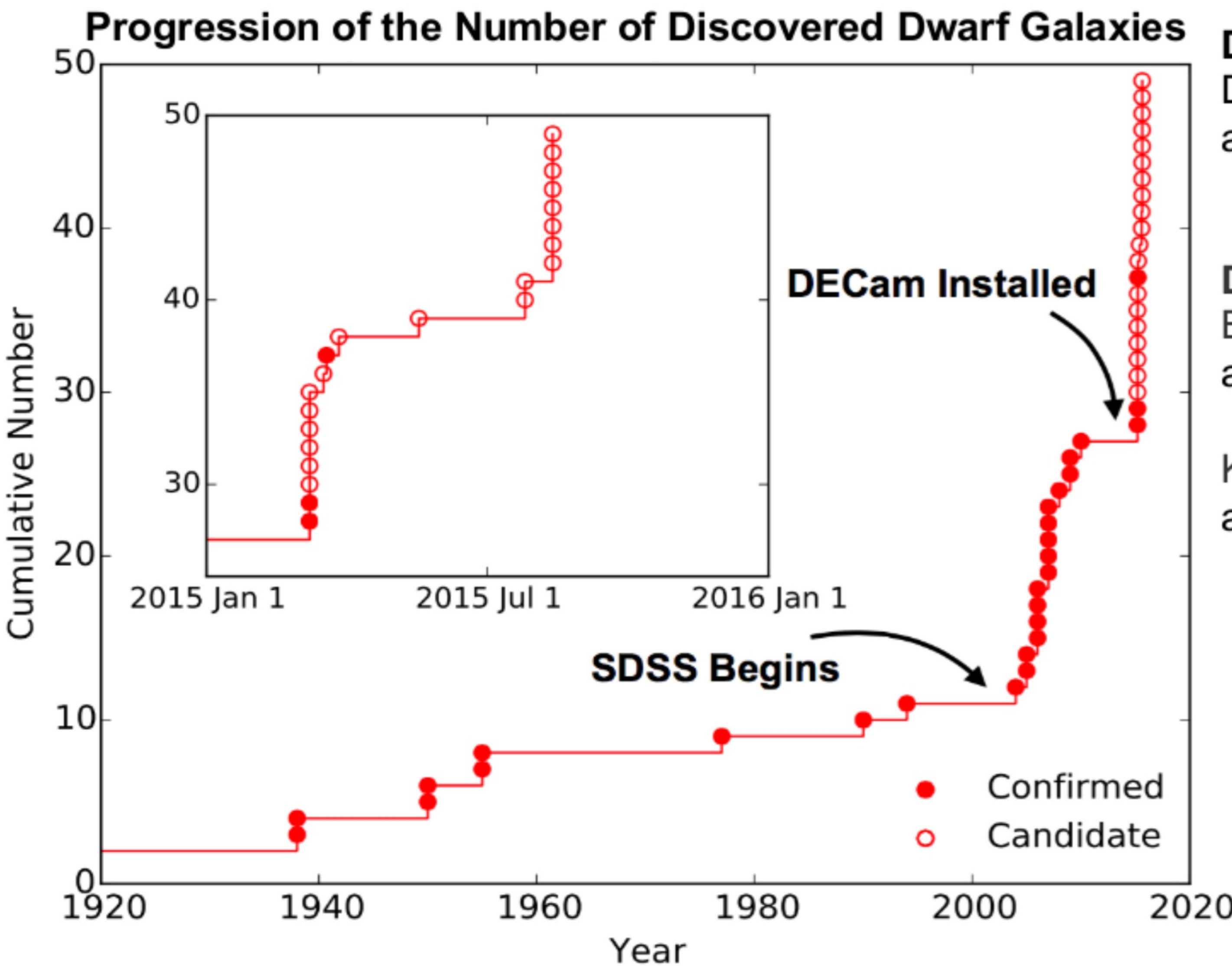
Classical Dwarf spheroidal galaxies: promising targets for DM detection



Dark Matter in the Milky Way (from simulations)



Dwarf Spheroidal Galaxies: Growing number of known targets

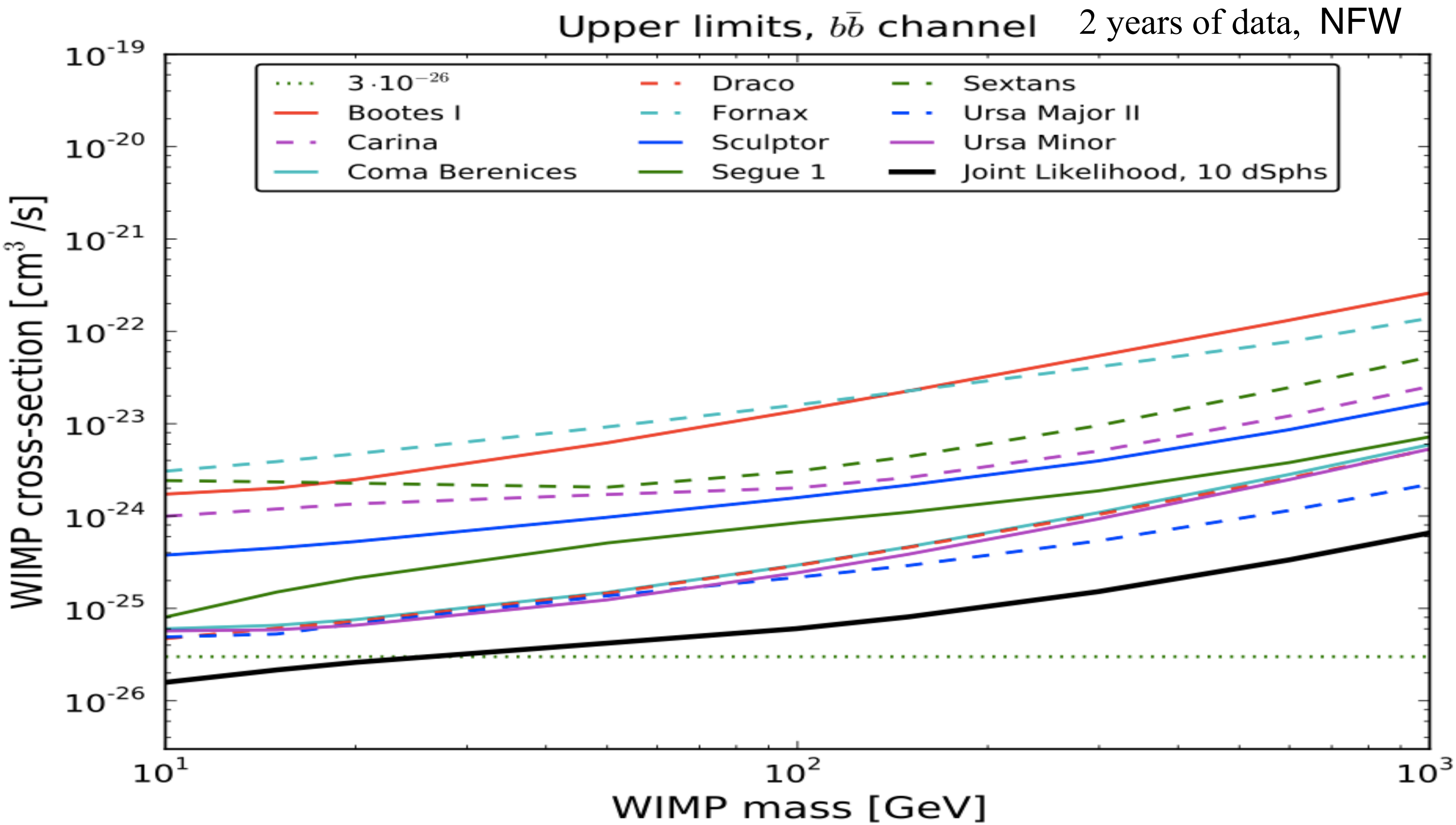


DES Year 2 Data: Drlica-Wagner+, arXiv:1508.03622

DES Year 1 Data: Bechtol+: arXiv:1503.02584

Koposov+: arXiv:1503.02079

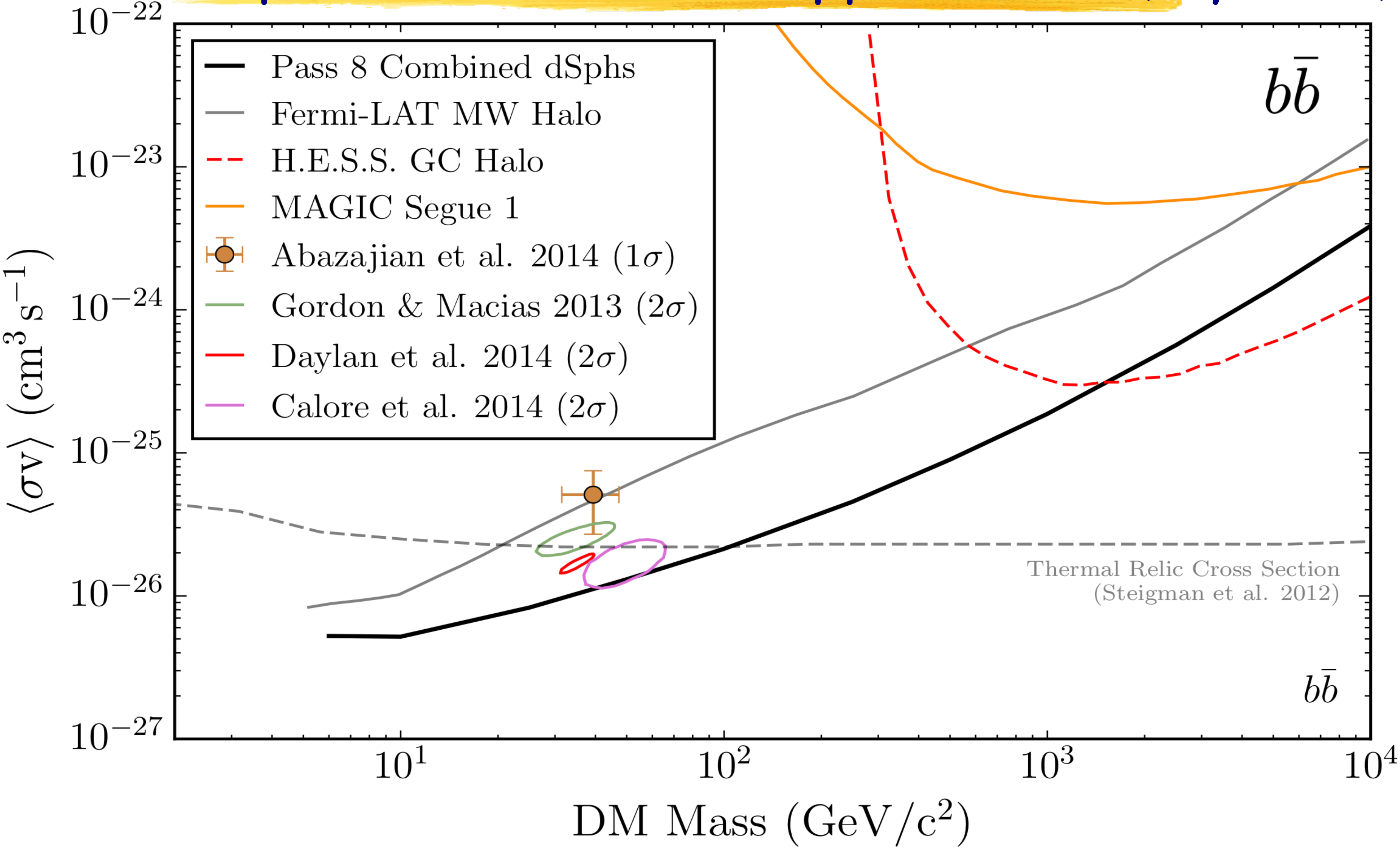
Dwarf Spheroidal Galaxies combined analysis



robust constraints including J-factor uncertainties from the stellar data statistical analysis

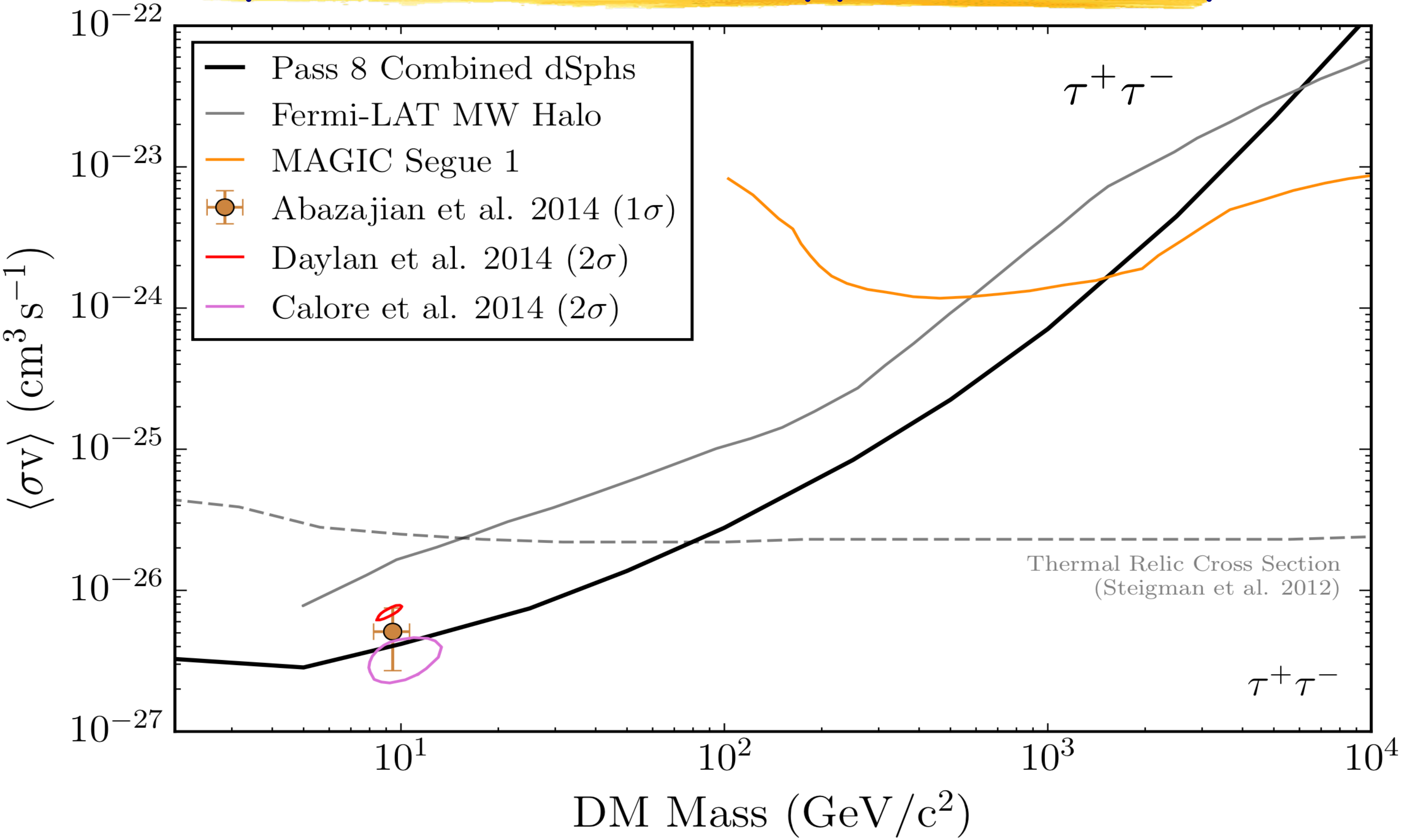
Fermi Lat Coll., PRL 107, 241302 (2011) [arXiv:1108.3546]

Dwarf Spheroidal Galaxies upper-limits (6 years)



M.Ackermann et al., [Fermi Coll.] PRL 115, 231301 (2015) [arXiv:1503.02641]

Dwarf Spheroidal Galaxies upper-limits (6 years)



M.Ackermann et al., [Fermi Coll.] PRL 115, 231301 (2015) [arXiv:1503.02641]

Combining all dSph observations

• Combination of the observation results towards 20 dwarf spheroidal galaxies (dSphs)

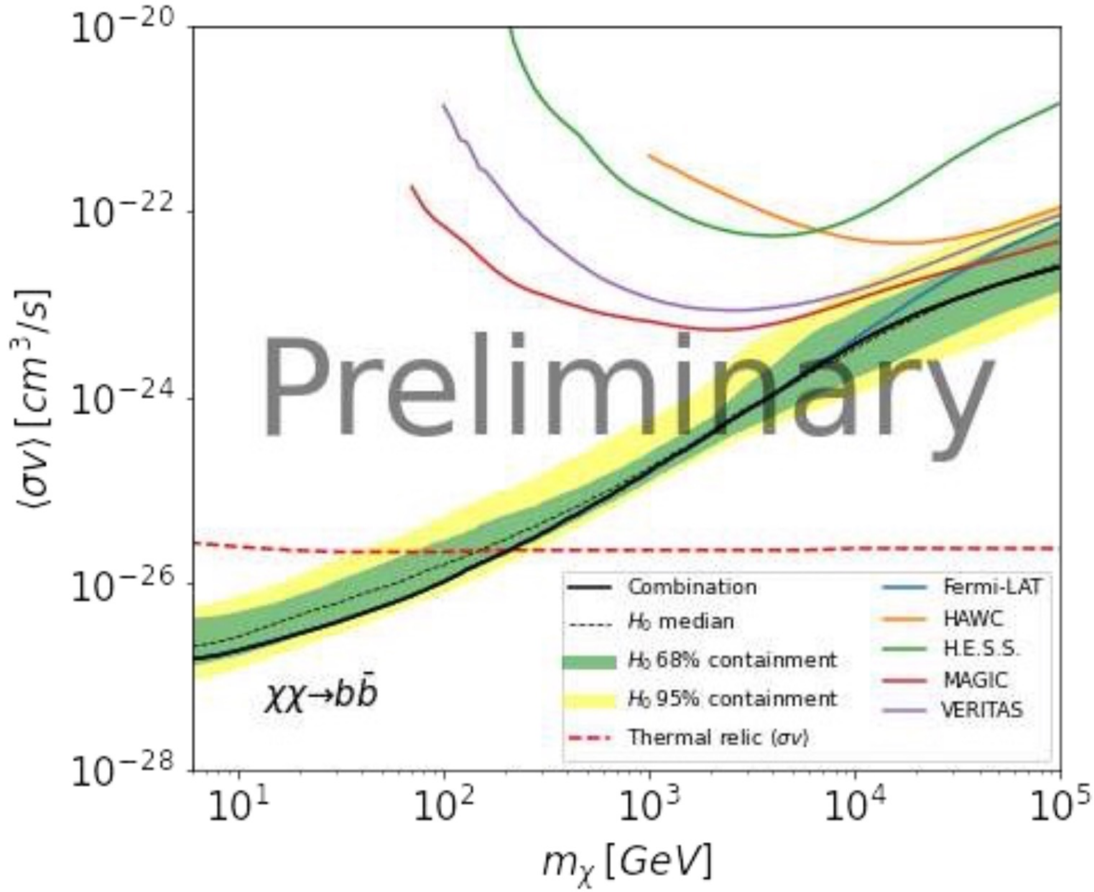
- Significant increase of the statistics

-> Increase the sensitivity to potential dark matter signals

- Cover the widest energy range ever investigated : 20 MeV – 80 TeV

• Common elements :

- Agreed model parameters
- Sharable likelihood table formats
- Joint likelihood test statistic

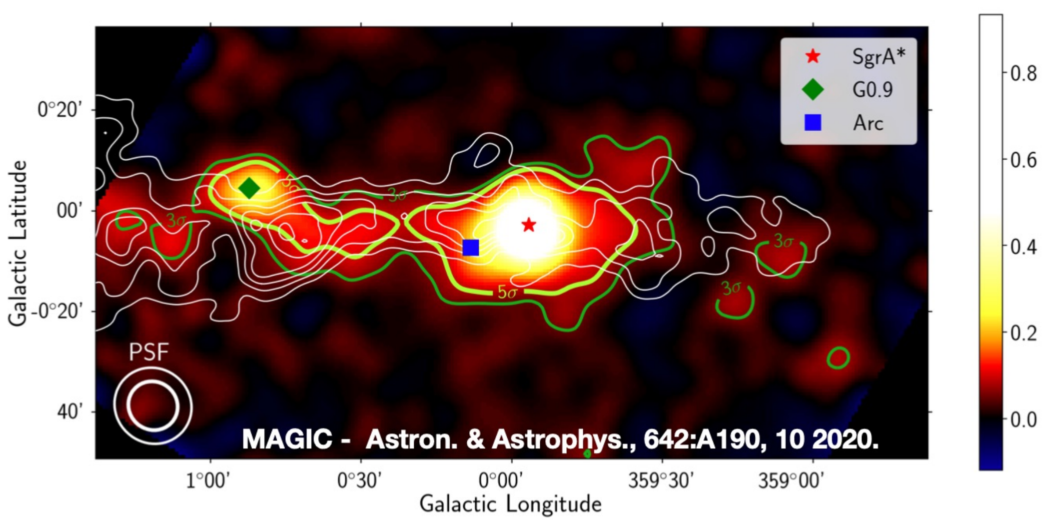


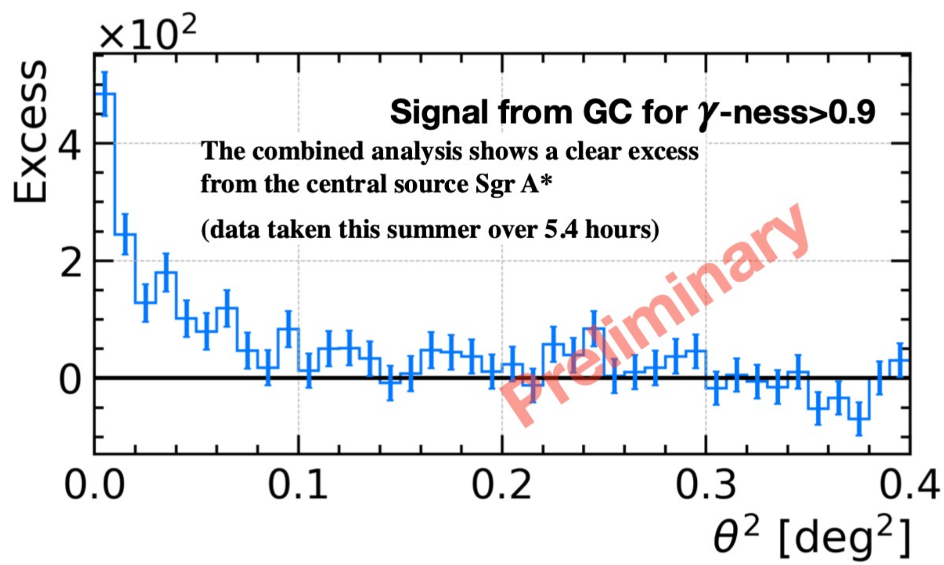
Joint MAGIC –LST 1 observation of Galactic Center

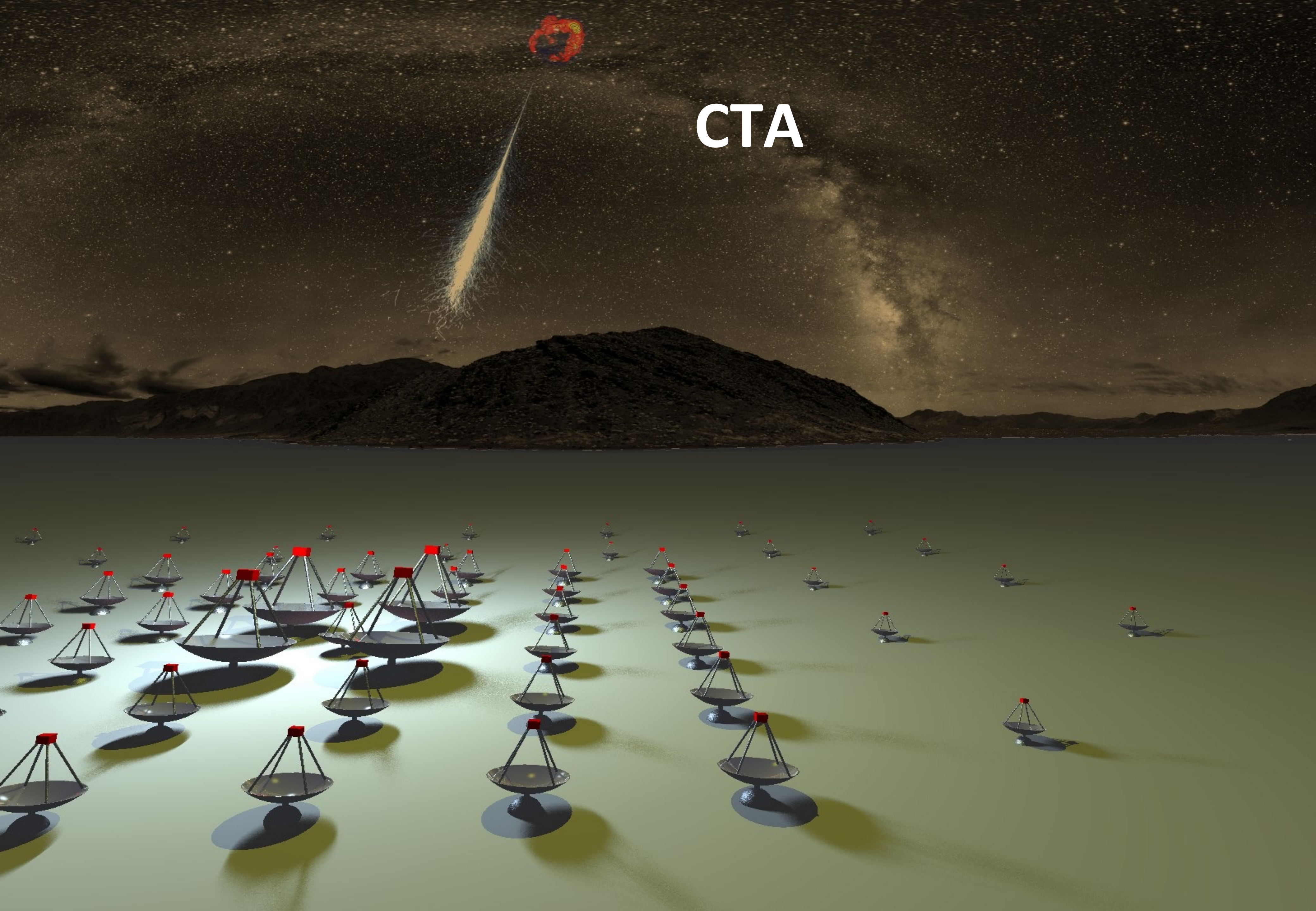
•The Galactic Center (GC) is one of the target regions for MAGIC + LST-1 observations, given the abundance of science targets (Sgr A*, Gal. diffuse emission and Dark Matter)

•The GC region culminates at large zenith angle of 58 degree seen at La Palma, thus enlarging the light pool and increasing the efficiency of the stereo triggering of LST array

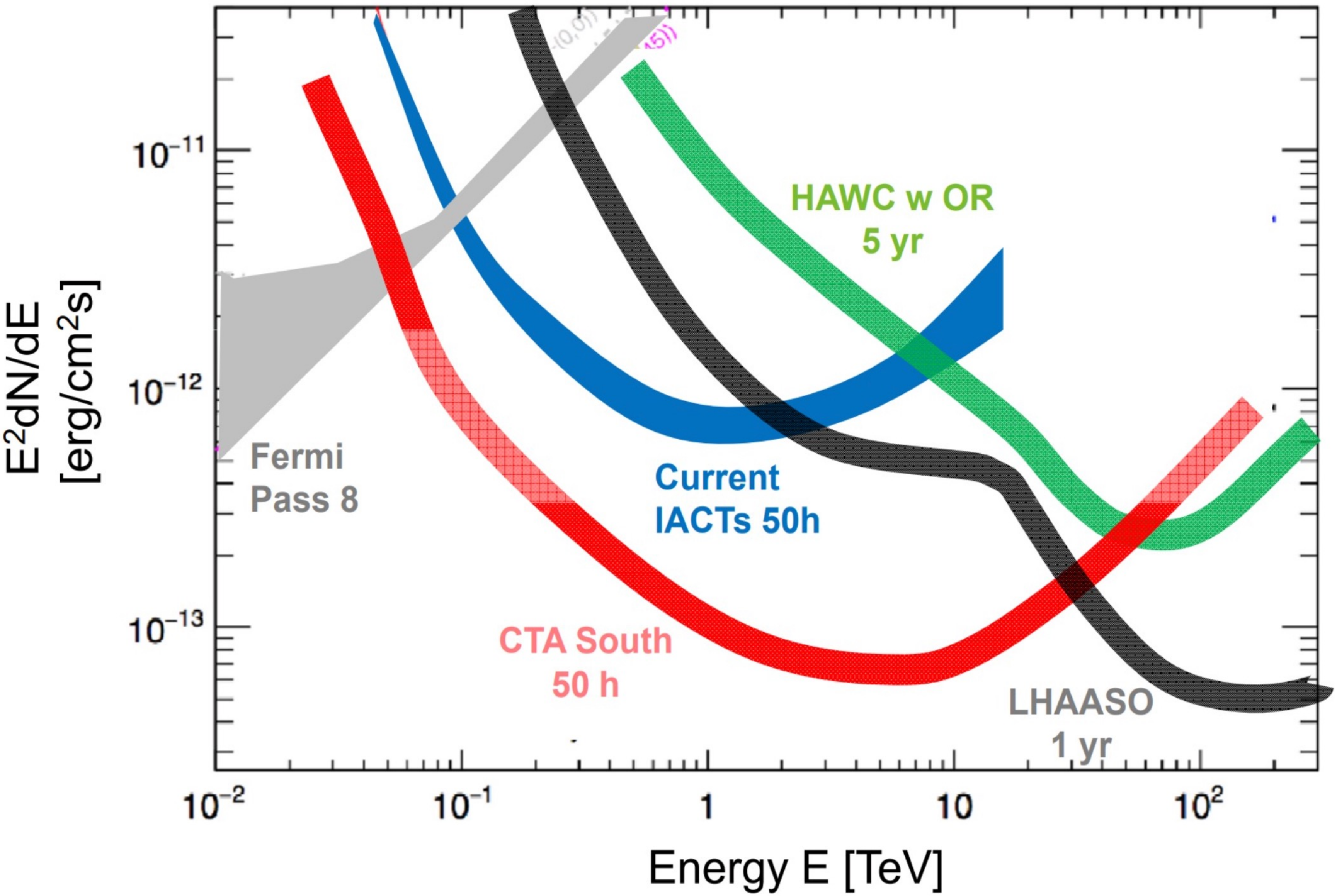
•At the same time, the complexity of the area, requires improved the angular resolution in order to understand/constrain the origin of the gamma-ray emission



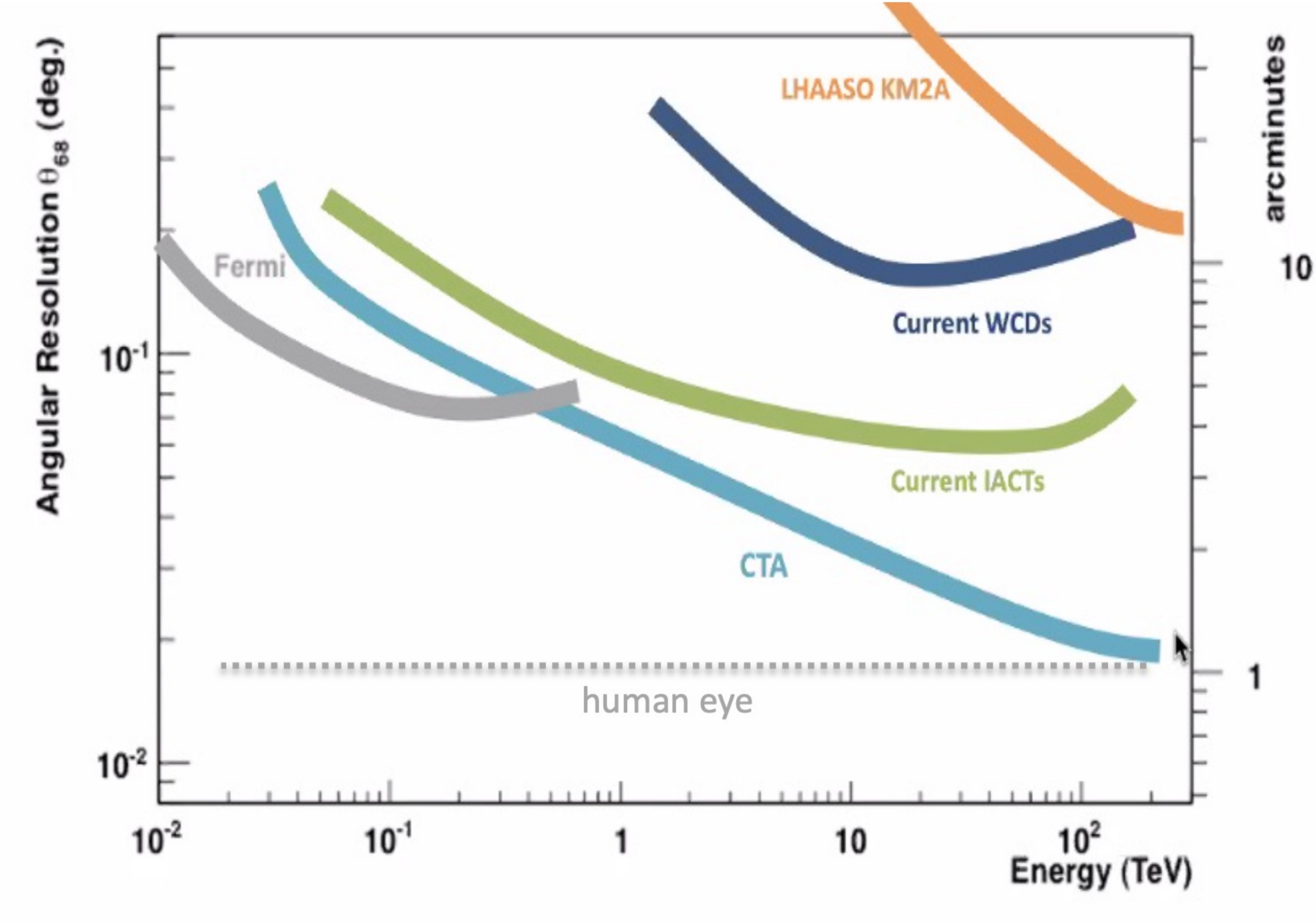



CTA

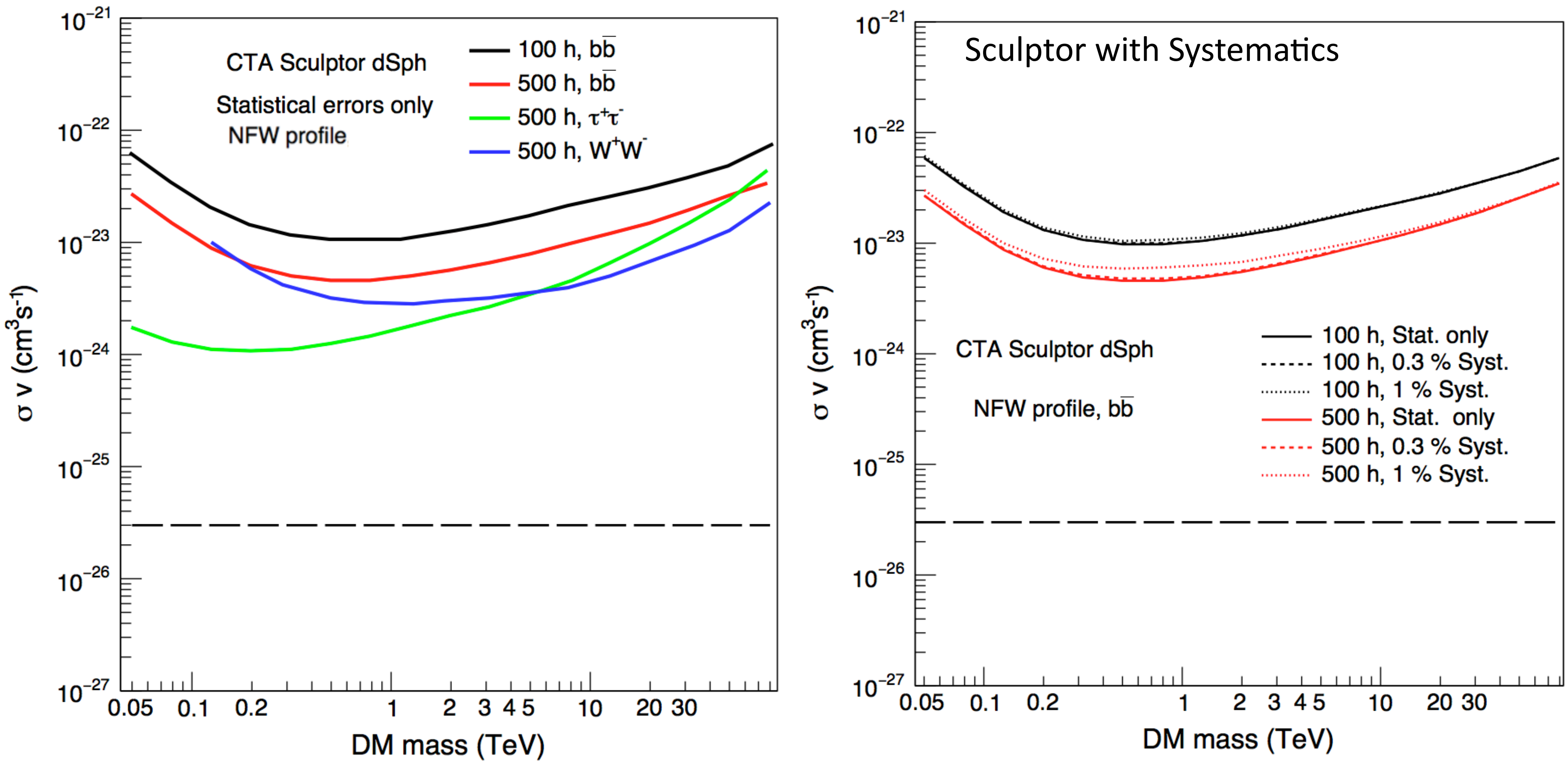
γ-ray detectors sensitivities



Angular resolution



Dwarf Spheroidal Galaxies: CTA Sensitivity



There are several of the newly discovered dSph that have a better case for being a promising target,

Will choose most promising targets before observations with the latest knowledge.

A.M. PoS(ICRC2017)921

Measuring DM densities in dSph halos

Optimal dSphs selected according to:
1. Distance(d<100pc)
2. Culmination zenith angle (ZAmin < 40°)

Targets with no/poor brightness and/ or kinematic data excluded from the MCMC Jeans analysis.

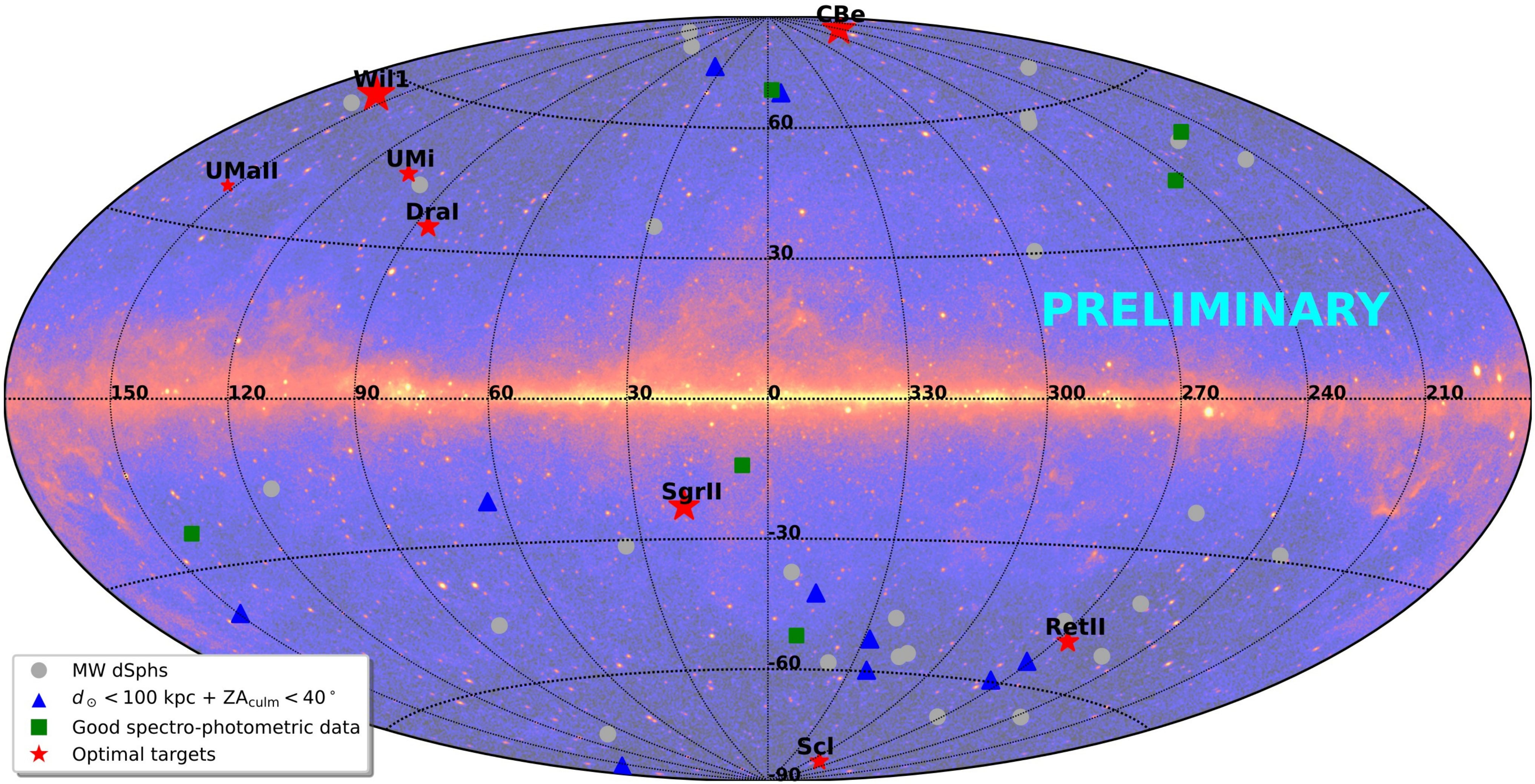
Surviving sample:

— 6 Northern dSphs (1 classical + 5 ultra-faint)

— 6 Southern dSphs (3 classical + 3 ultra-faint)

Name	Abbr.	Type	R.A. (hh mm ss)	dec. (dd mm ss)	Distance (kpc)	ZA_{culm} N (deg)	ZA_{culm} S (deg)	Month
Andromeda XVIII	AndXVIII	uft	00 02 14.5	+45 05 20	1330 ± 104	16.3	69.7	Sep
Aquarius	Aqr	uft	20 46 51.8	−12 50 53	1030 ± 57	41.6	11.8	Aug
Boötes I	BoöI	uft	14 00 06.0	+14 30 00	65 ± 3	14.3	39.1	Apr
Boötes II	BoöII	uft	13 58 00.0	+12 51 00	39 ± 2	15.9	37.5	Apr
Boötes III	BoöIII	uft	13 57 12.0	+26 48 00	46 ± 2	2.0	51.4	Apr
Canes Venatici I	CVnI	uft	13 28 03.5	+33 33 21	216 ± 8	4.8	58.2	Apr
Canes Venatici II	CVnII	uft	12 57 10.0	+34 19 15	159 ± 8	5.6	58.9	Apr
Carina	Car	cls	06 41 36.7	−50 57 58	106 ± 1	79.7	26.3	Dec
Cetus I	CetI	uft	00 26 11.0	−11 02 40	748 ± 31	39.8	13.6	Sep
Cetus II	CetII	uft	01 17 52.8	−17 25 12	30 ± 3	46.2	7.2	Oct
Columba I	ColI	uft	05 31 26.4	−28 01 48	182 ± 18	56.8	3.4	Dec
Coma Berenices	CBe	uft	12 26 59.0	+23 54 15	42 ± 2	4.9	48.5	Mar
Draco I	DraI	cls	17 20 12.4	+57 54 55	75 ± 4	29.2	82.5	Jun
Draco II	DraII	uft	15 52 47.6	+64 33 55	20 ± 3	35.8	89.2	May
Eridanus II	EriII	uft	03 44 21.5	−43 31 48	330 ± 16	72.3	18.9	Nov
Eridanus III	EriIII	uft	02 22 45.5	−52 16 48	95 ± 27	81.0	27.7	Oct
Fornax	For	cls	02 39 59.3	−34 26 57	146 ± 1	63.2	[illegible]	Oct
Grus I	GruI	uft	22 56 42.4	−50 09 48	120 ± 17	[illegible]	[illegible]	Sep
Grus II	GruII	uft	22 04 04.8	−46 26 24	53 ± 5	75.2	21.8	Aug
Hercules	Her	uft	16 31 02.0	+12 47 30	137 ± 11	[illegible]	37.4	May
Horologium I	HorI	uft	02 55 28.9	−54 06 36	[illegible]	82.9	29.5	Oct
Hydra II	HyaII	uft	12 21 42.1	−31 59 07	[illegible]	60.7	7.4	Mar
Indus I	IndI	uft	21 08 48.1	[illegible]	69 ± 16	79.9	26.5	Aug
Indus II	IndII	uft	20 38 52.8	[illegible]	214 ± 16	74.9	21.5	Aug
Laevens 3	Lae3	uft	[illegible]	+14 58 48	67 ± 3	13.8	39.6	Aug
Leo I	LeoI	cls	[illegible]	+12 18 23	272 ± 10	16.5	36.9	Feb
Leo II	LeoII	cls	11 13 28.8	+22 09 06	240 ± 9	6.6	46.8	Mar
Leo IV	LeoIV	uft	11 32 57.0	−00 32 00	151 ± 4	29.3	24.1	Mar
Leo V	LeoV	uft	11 31 09.6	+02 13 12	169 ± 5	26.5	26.9	Mar
Leo T	LeoT	uft	09 34 53.4	+17 03 05	377 ± 28	11.7	41.7	Feb
Phoenix I	PheI	uft	01 51 06.3	−44 26 41	427 ± 31	73.2	19.8	Oct
Phoenix II	PheII	uft	23 39 57.6	−54 24 36	95 ± 18	83.2	29.8	Sep
Pictor I	PicI	uft	04 43 48.0	−50 16 48	126 ± 24	79.0	25.7	Nov
Pisces II	PscII	uft	22 58 31.0	+05 57 09	182 ± 13	22.8	30.6	Sep
Reticulum II	RetII	uft	03 35 40.9	−54 03 00	32 ± 2	82.8	29.4	Nov
Reticulum III	RetIII	uft	03 45 26.3	−60 27 00	92 ± 13	89.2	35.8	Nov
Sagittarius I	SgrI	dis	18 55 19.5	−30 32 43	31 ± 1	59.3	5.9	Jul
Sagittarius II	SgrII	uft	19 52 40.5	−22 04 05	67 ± 5	50.8	2.6	Jul
Sculptor	Scl	cls	01 00 09.4	−33 42 33	84 ± 2	62.5	9.1	Oct
Segue 1	Seg1	uft	10 07 04.0	+16 04 55	23 ± 2	12.7	40.7	Feb
Segue 2	Seg2	uft	02 19 16.0	+20 10 31	36 ± 2	8.6	44.8	Oct
Sextans	Sex	cls	10 13 03.0	−01 36 53	84 ± 3	30.4	23.0	Feb
Triangulum II	TriII	uft	02 13 17.4	+36 10 42	30 ± 2	7.4	60.8	Oct
Tucana I	TucI	uft	22 41 49.6	−64 25 10	855 ± 35	—	39.8	Sep
Tucana II	TucII	uft	22 52 16.7	−58 33 36	58 ± 6	87.3	33.9	Sep
Tucana III	TucIII	uft	23 56 35.9	−59 36 00	25 ± 2	88.4	35.0	Sep
Tucana IV	TucIV	uft	00 02 55.3	−60 51 00	48 ± 4	89.6	36.2	Sep
Ursa Major I	UMaI	uft	10 34 52.8	+51 55 12	105 ± 2	23.2	76.6	Mar
Ursa Major II	UMaII	uft	08 51 30.0	+63 07 48	35 ± 2	34.4	87.8	Feb
Ursa Minor	UMi	cls	15 09 08.5	+67 13 21	68 ± 2	38.5	—	May
Willman 1	Wil1	uft	10 49 21.0	+51 03 00	38 ± 7	22.3	75.7	Mar

Dwarf Spheroidal Galaxies: Selection of optimal candidates for CTA



Optimal dSphs selected according to:

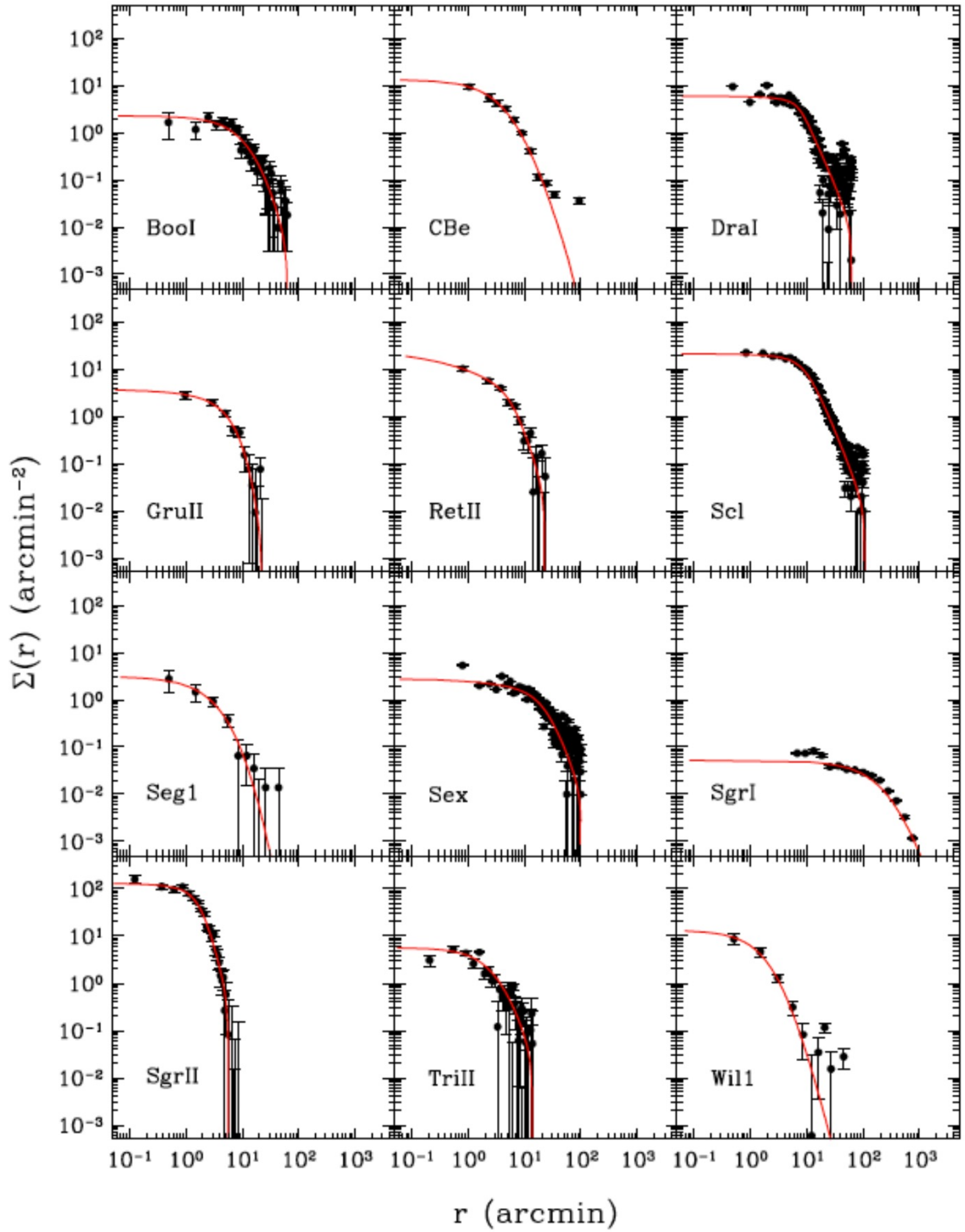
1.distance (d < 100 pc) 2. culmination zenith angle (ZA < 40°) 3. availability of good spectro-photometric data.

Surviving sample: 8 Northern dSphs (2 classical + 6 ultra-faint) 6 Southern dSphs (3 classical + 3 ultra-faint)

CTA Consortium in prep. 2023

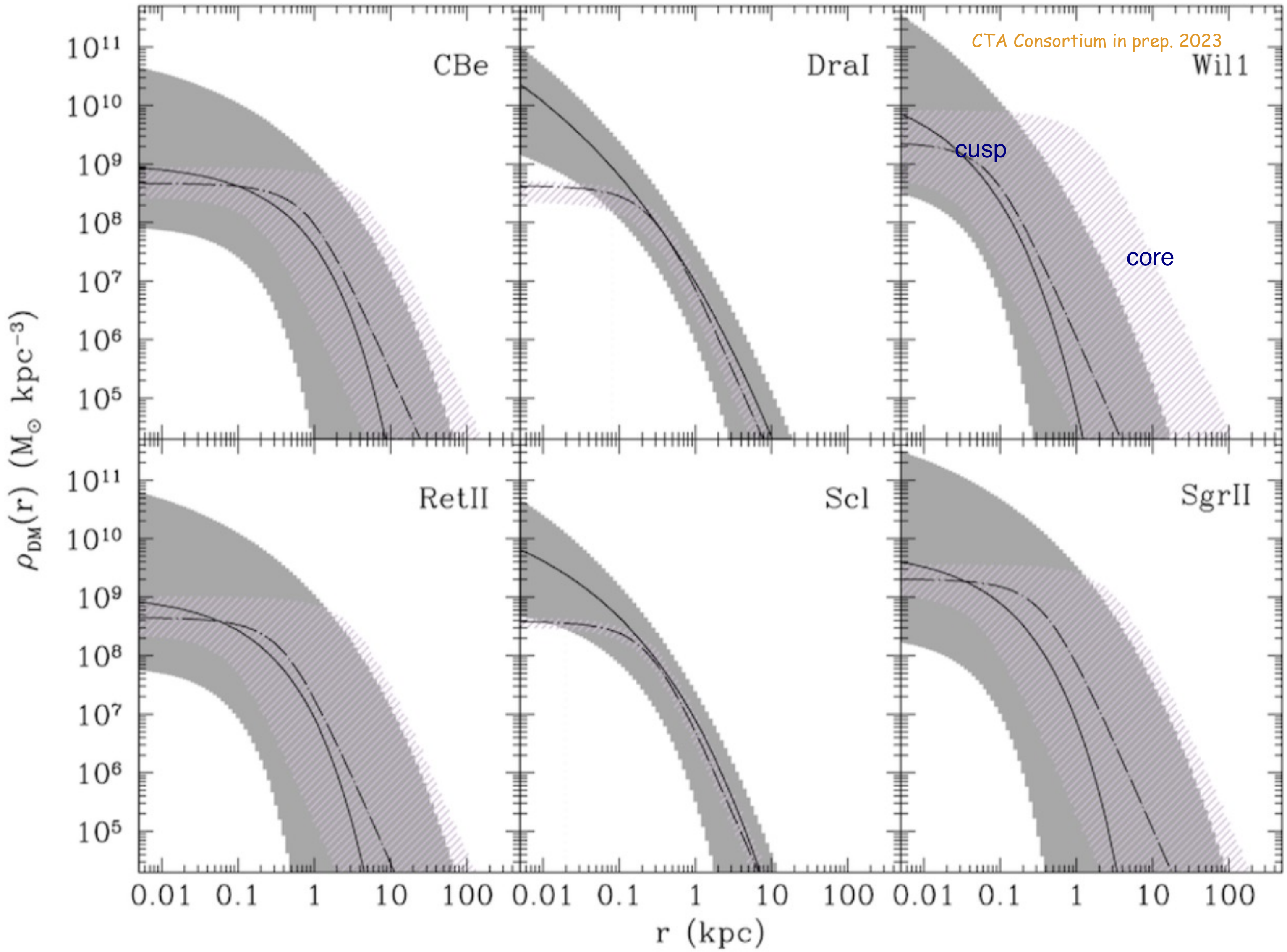
Measuring Dark Matter densities in dSph halos

Best-fit brightness profiles Σ(r) of the analyzed dSphs as a function of the object's projected (2D) radial coordinate r from the dSph centroid



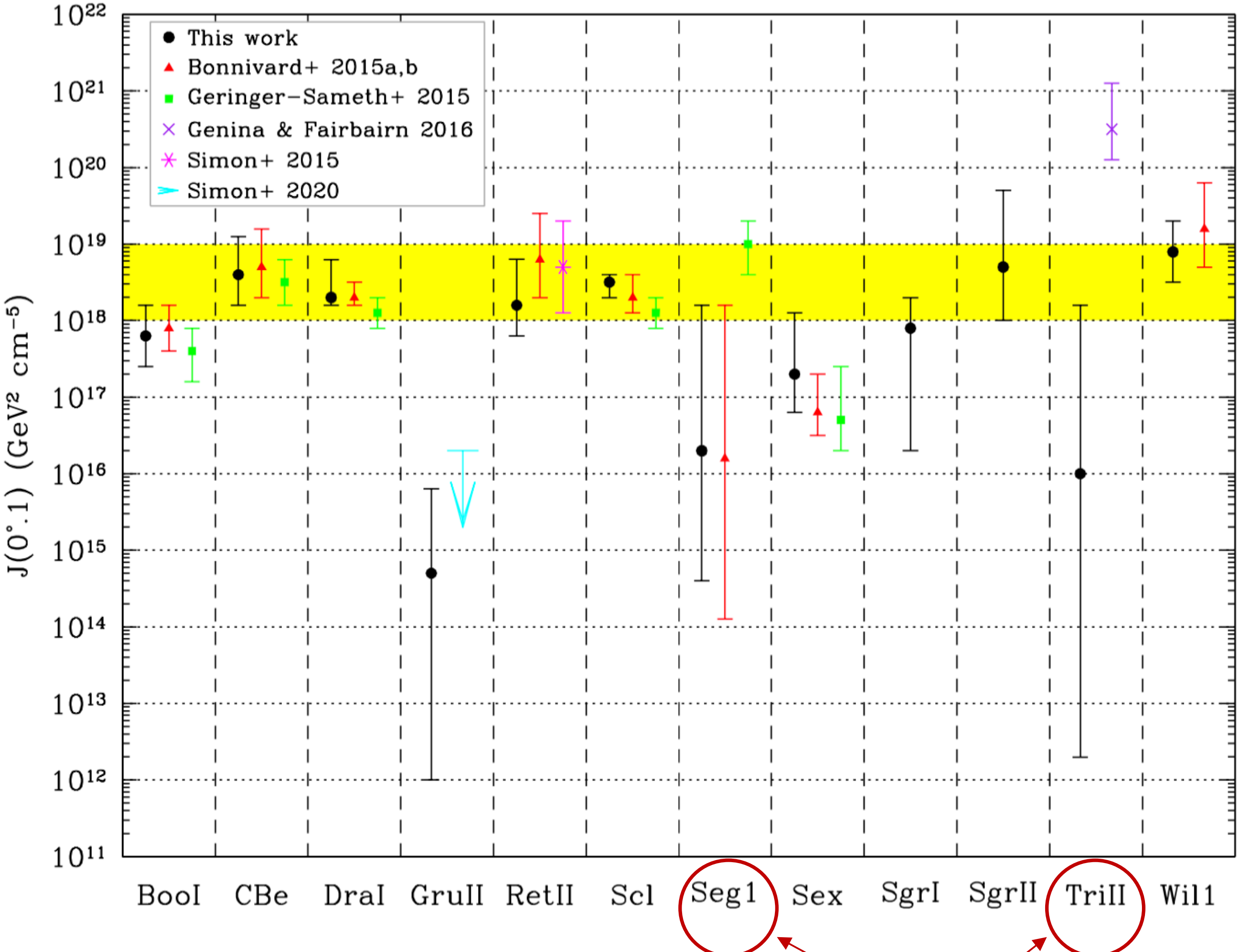
Measuring Dark Matter densities in dSph halos

Core and cusp DM density profiles for 6 dSphs targets



Measuring Dark Matter densities in dSph halos

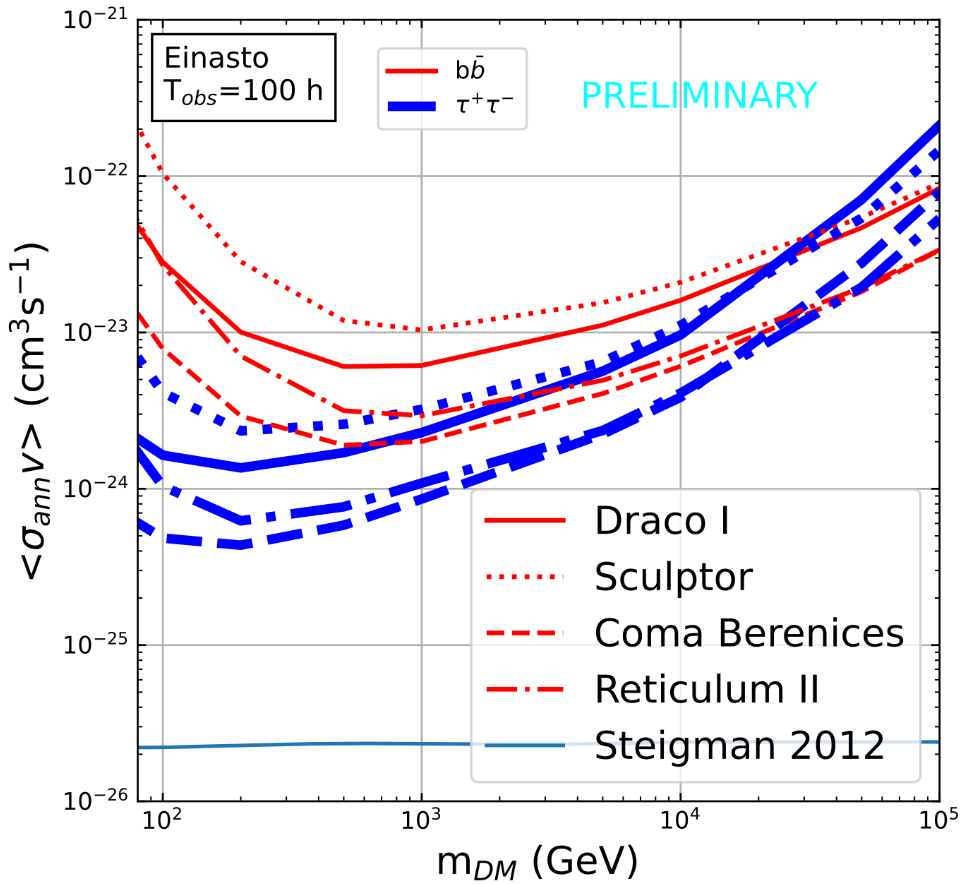
Comparison of astrophysical factor for DM annihilation within 0.1 deg of integration



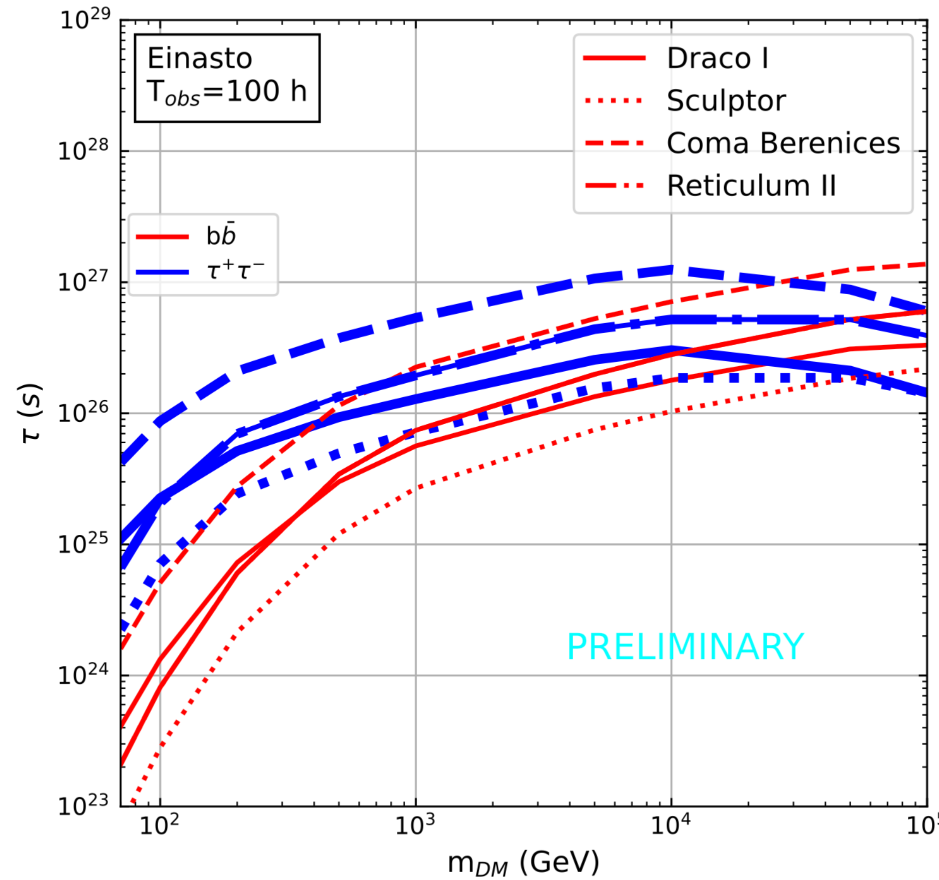
CTA Consortium paper in prep. and ICRC 2023 by the dSph task force F. G. Saturni, M. Doro, A. Morselli, G. Rodríguez-Fernández

Dwarf Spheroidal Galaxies: CTA Sensitivity

upper limits on annihilating DM cross sections



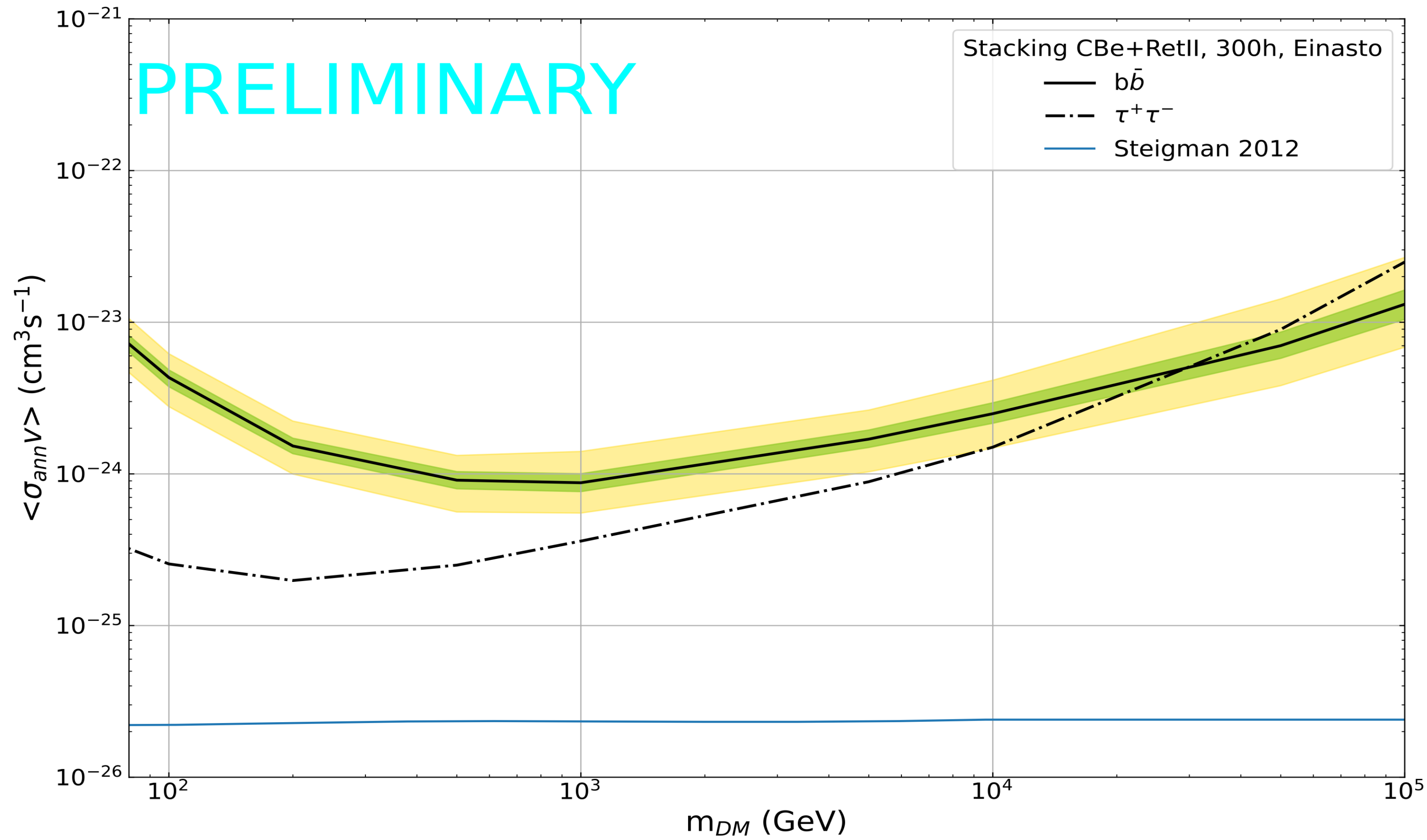
lower limits on the particle lifetime



computed with the Einasto DM density profile derived by CLUMPY, computed assuming 100 h of observation for annihilation in the two pure DM channels $b\bar{b}$ and $\tau^+\tau^-$

CTA Consortium paper in prep. and ICRC 2023 by the dSph task force F. G. Saturni, M. Doro, A. Morselli, G. Rodríguez-Fernández

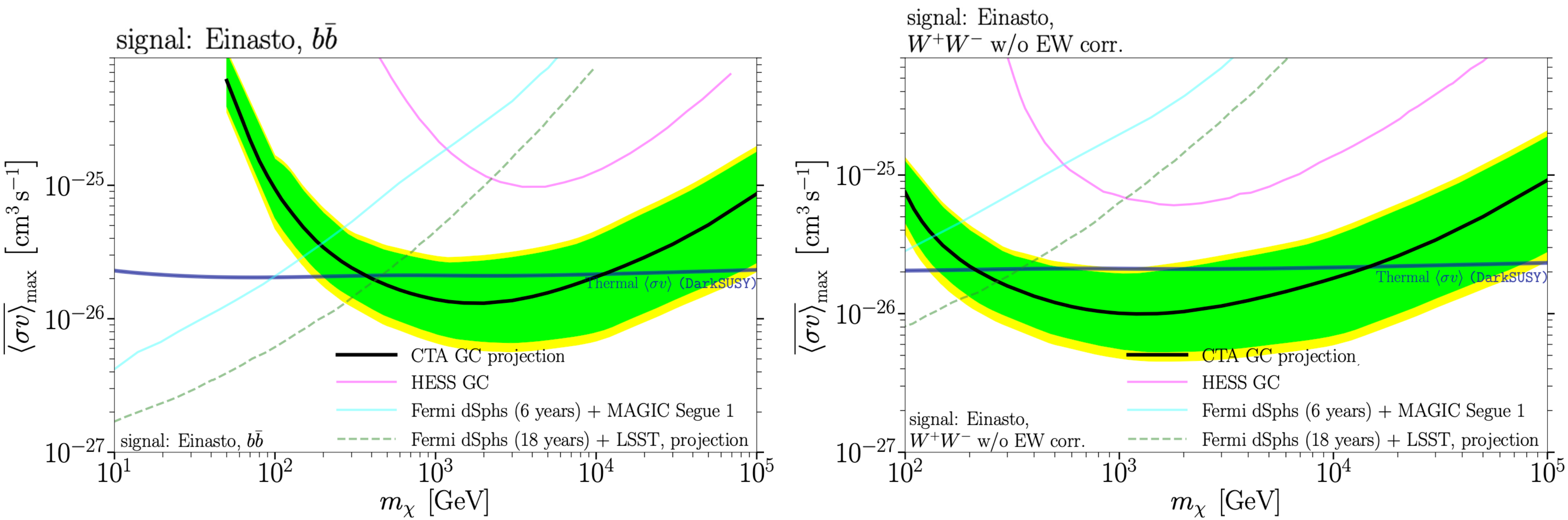
Dwarf Spheroidal Galaxies: Stacking analysis



- Stacking analysis performed on 600 h of CBe+RetII observations (300 h each) for the DM annihilation into $b\bar{b}$ and $\tau^+\tau^-$
- Depending on the choice of target and the interaction channel, CTA can reach x-sections <10- 24 cm3 s-1 for DM particle masses >1 TeV.

CTA Consortium paper in prep. and ICRC 2023 by the dSph task force F. G. Saturni, M. Doro, A. Morselli, G. Rodríguez-Fernández

Galactic center CTA Sensitivity



- Einasto profile

525 h

$$\rho_{DM} = \rho_s \exp\left[-\frac{\alpha}{2}\left(\frac{r}{r_s}\right)^\alpha - 1\right], \; J \sim 7.1 \times 10^{22} \mathrm{GeV}^2/\mathrm{cm}^5$$

- Main source of background : sources, Fermi Bubble, interstellar γ, residual CR

The CTA Consortium JCAP01(2021) 057 January 27, 2021 [arXiv:2007.16129]

CTA Key Science Project Targets

- Galactic Center

high DM density but high astrophysical emissions

- dSph

no background but low signal

- LMC

neaby & massive but astrophysical emissions

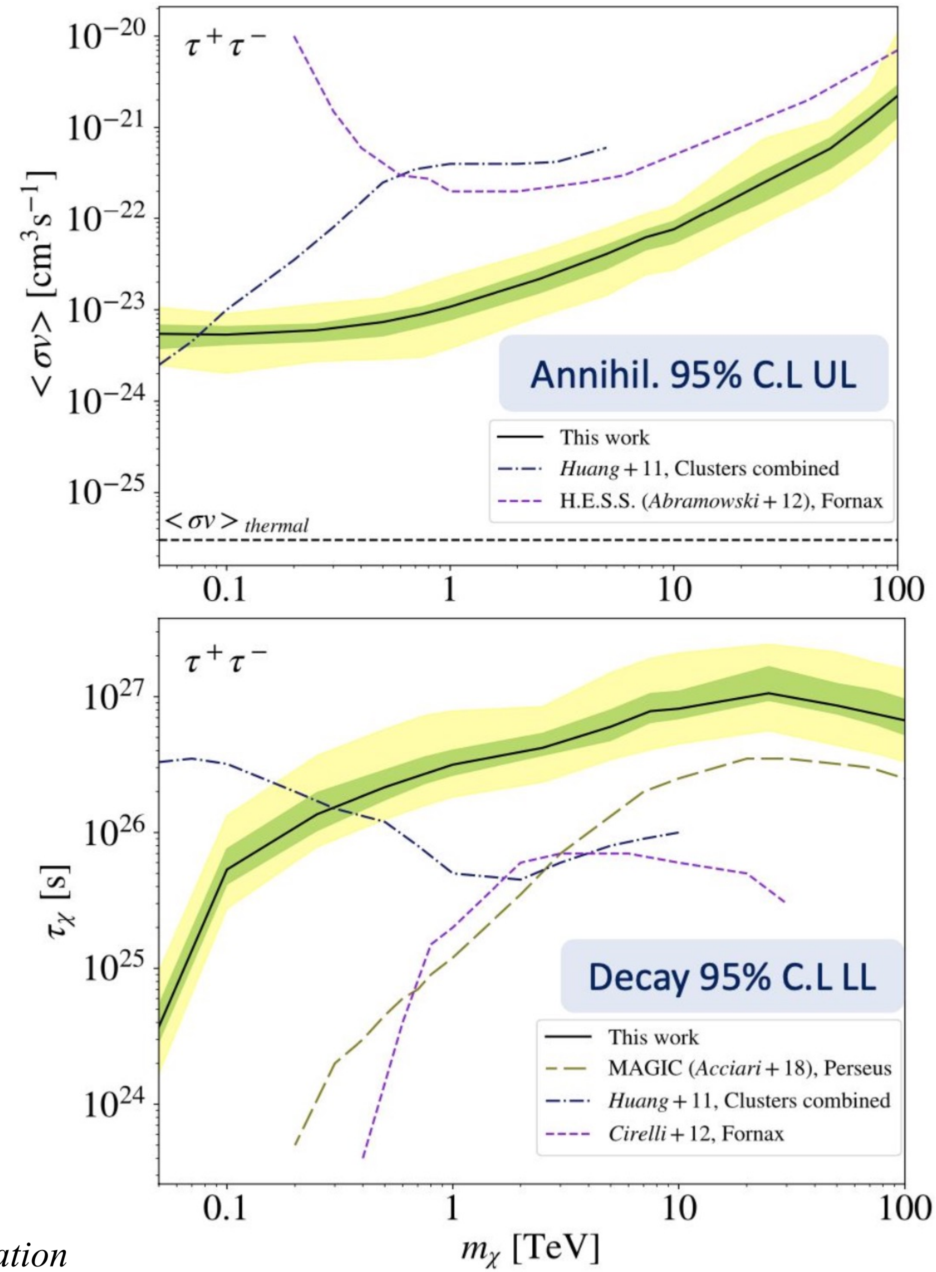
- galaxy cluster

very massive (best for decay)

Perseus cluster

Expected CTA sensitivity

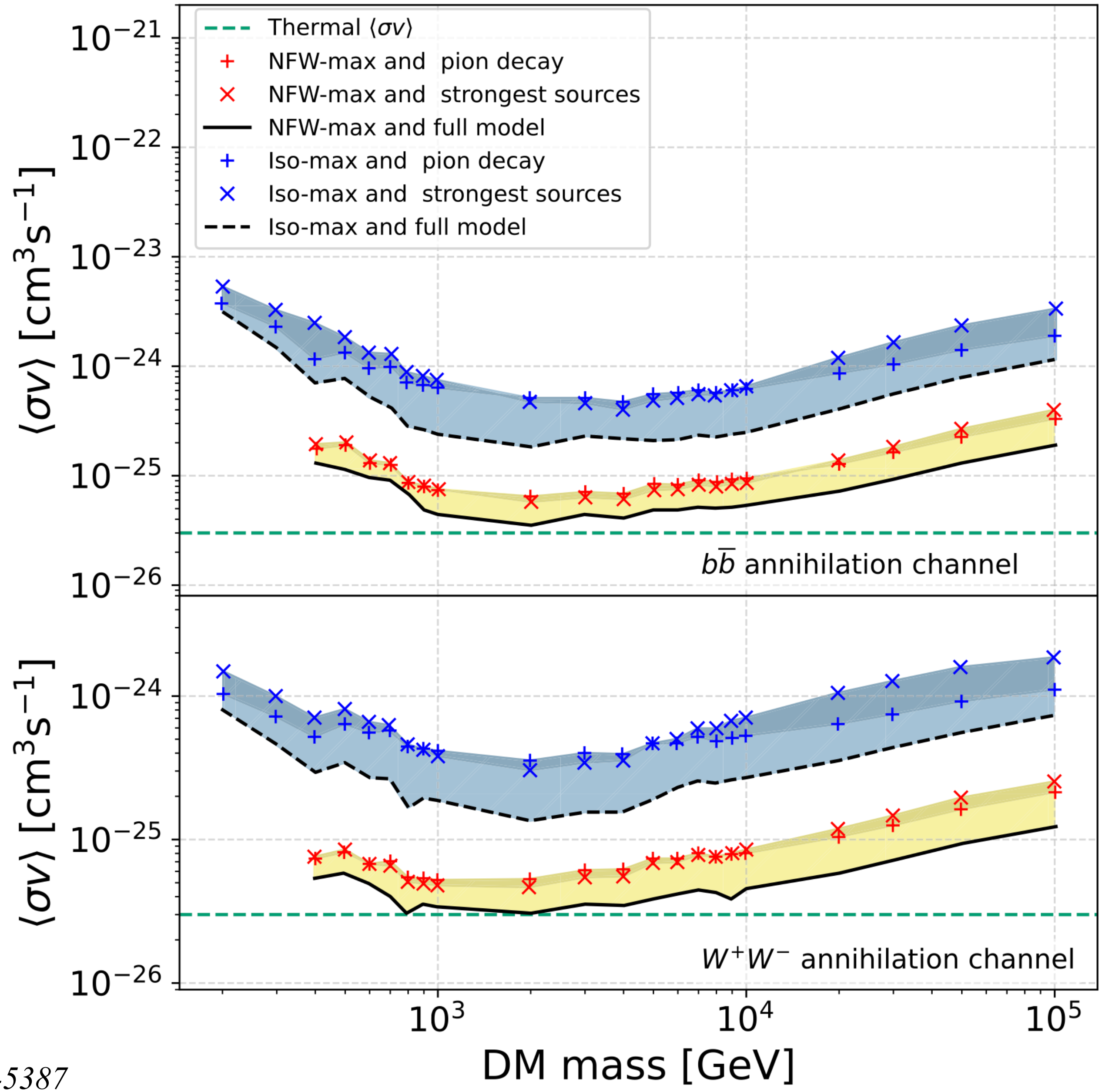
- Distance: 75 Mpc
- Dark matter content: $\log_{10}$J [GeV^{2}cm^{-5}] =18.43
  $\log_{10}$D [GeV^{2}cm^{-2}]=19.20
- Observation time: 300 h
- Astrophysical gamma-ray sources:
  - Active Galactic Nuclei
  - diffuse emission (CR interactions)



The CTA Consortium, in preparation

DM searches in the Large Magellanic Cloud

- Distance: 50.1 kpc
- Dark matter content: $\log_{10} J[\mathrm{GeV}^2\mathrm{cm}^{-5}] = 21.14$
- Observation time: 340 h
- Astrophysical gamma-ray sources:
  - 4 known sources: SNR, PWN
  - diffuse emission (CR interactions)



The CTA Consortium, MNRAS 523 (2023) 4, 5353-5387

CTA Search for Dark Matter beyond WIMP

Axion Like Particle (ALP) search prospects

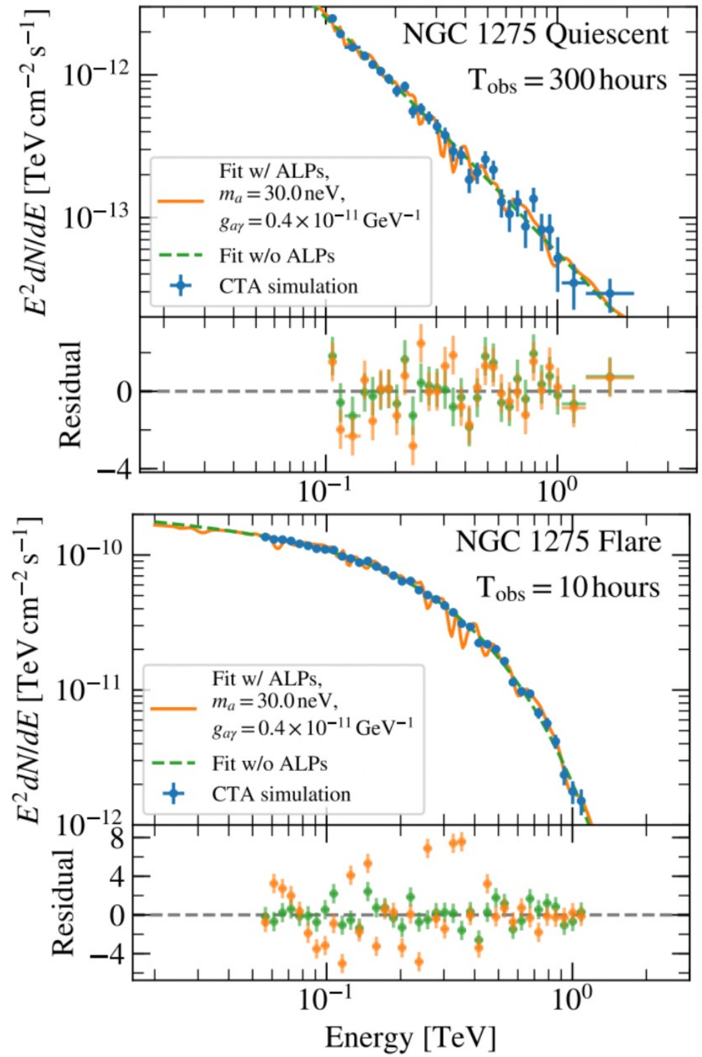
$\gamma + B \rightarrow a + B \rightarrow \gamma' + \ldots$

conversion probability ($E > E_{crit}$)

$$P_{a\gamma} \sim \sin^2\left(\frac{g_{a\gamma}Bl}{2}\right),$$

$$E_{\rm crit} \sim 2.5\ \mathrm{GeV} \times \left(\frac{|m_a - \omega_{\rm pl}|}{1\mathrm{neV}}\right)^2 \left(\frac{B}{1\mu\mathrm{G}}\right)^{-1} \left(\frac{g_{a\gamma}}{10^{-11}\mathrm{GeV}^{-1}}\right)^{-1}$$

Simulated spectra of the radio galaxy NGC 1275



the observation is simulated without an ALP effect and is modeled both without ALPs and with a fixed set of magnetic-field realization and ALP parameters that are excluded at 95 % confidence level by the flaring state simulation

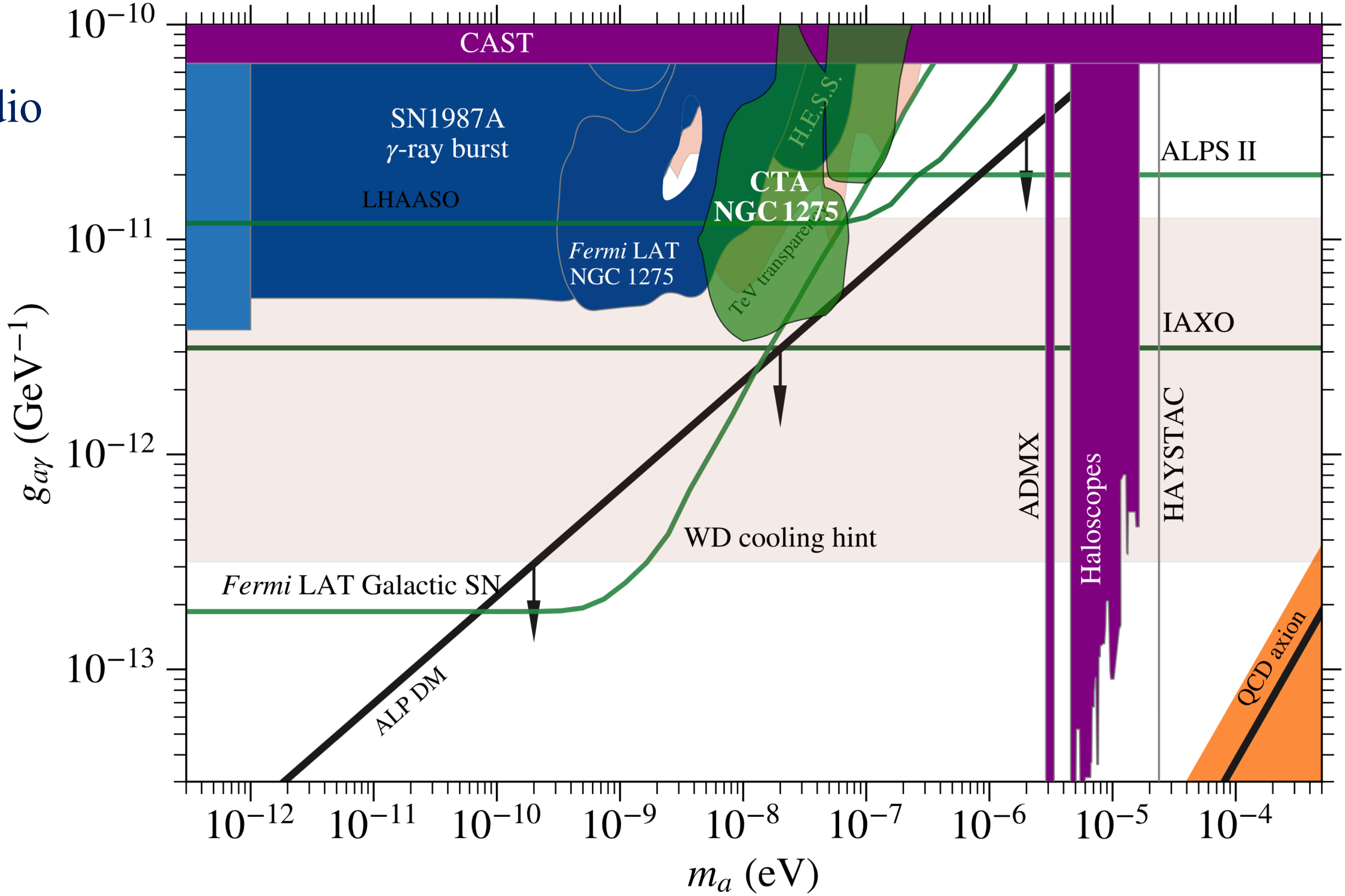
The CTA Consortium, JCAP 02 (2021) 048, 2021 [arXiv:2010.01349]

CTA Search for Dark Matter beyond WIMP

Axion Like Particle search prospects

- Observation of a flaring state of the radio galaxy NGC 1275 inside the Perseus cluster

- Observations of several AGN can be combined to further improve the CTA sensitivity.

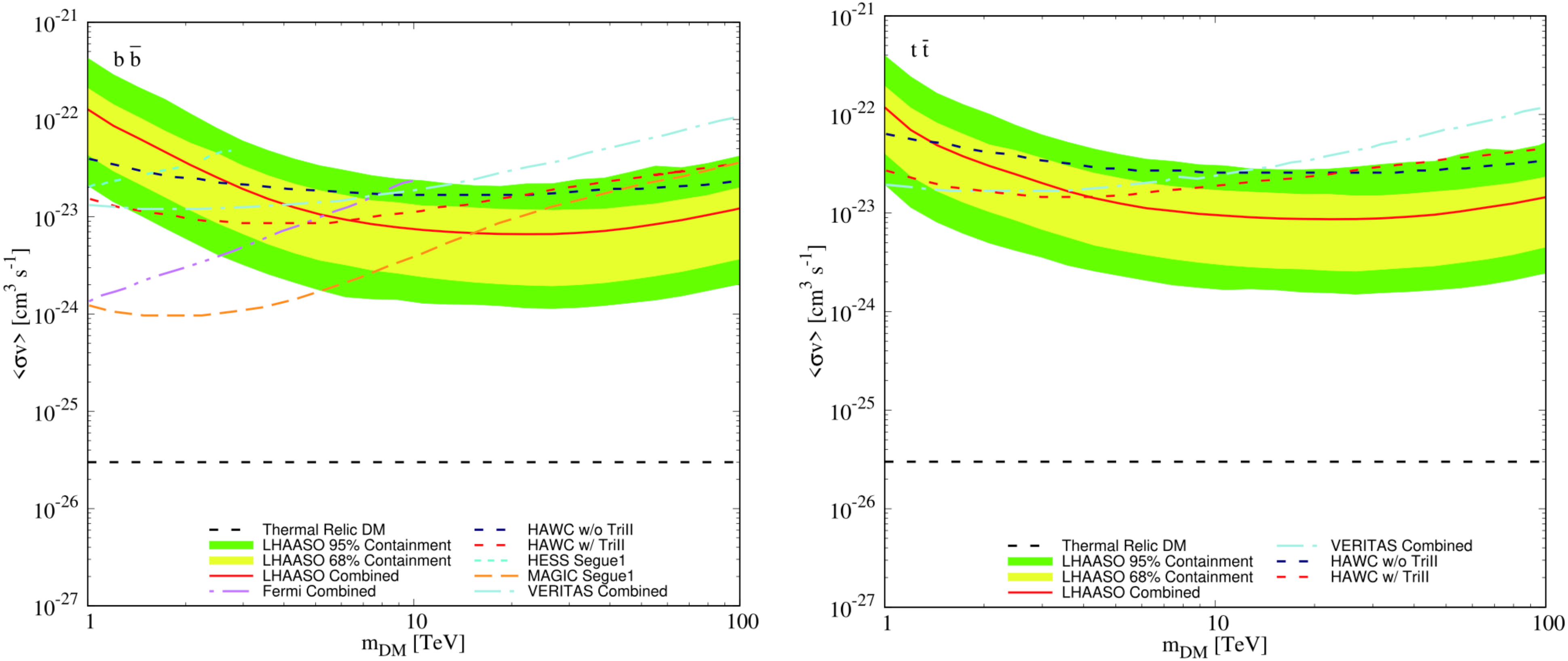


The CTA Consortium, JCAP 02 (2021) 048, 2021 [arXiv:2010.01349]

LHAASO

Mt. Haizi 4410 altitude

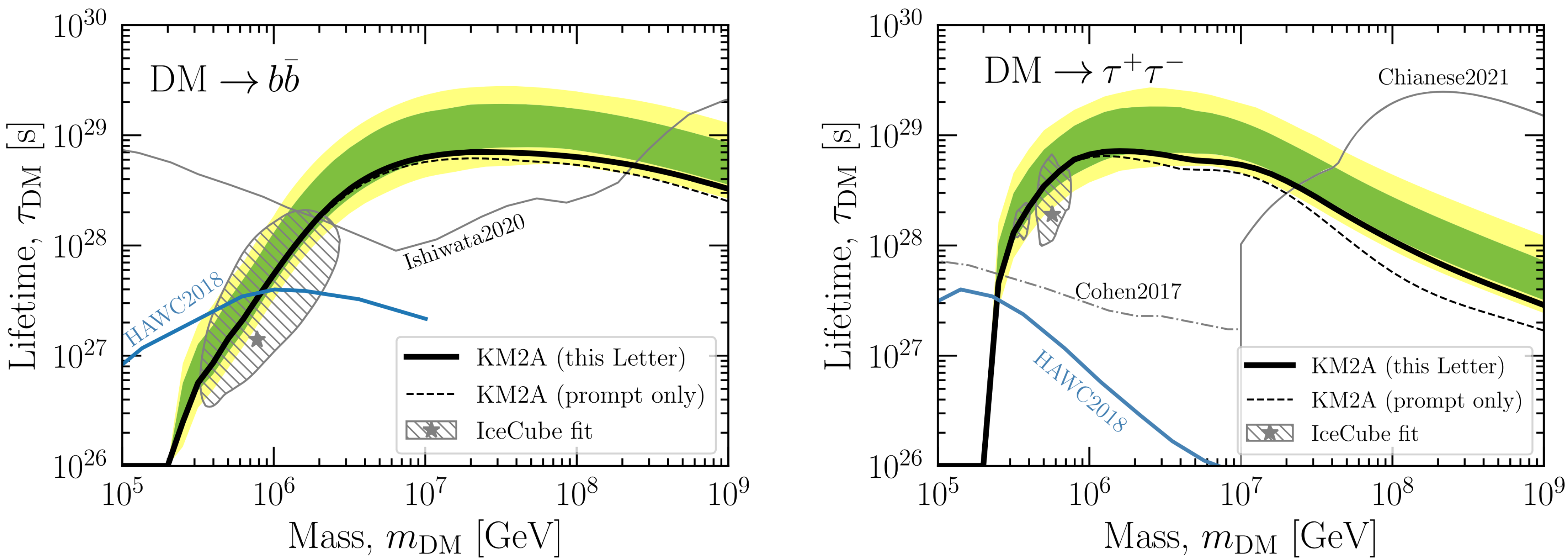
Combined one-year LHAASO sensitivities



The LHAASO median combined sensitivities (red solid lines) and related two-sided 68% (yellow bands) and 95% (green bands) containment bands of one year for the $b\bar{b}$, $\bar{t}t$ for 19 dSphs within the LHAASO FOV

Dong-Ze He et al., Phys. Rev. D 100, 083003 (2019)

Constraints on Heavy Decaying Dark Matter from 570 Days of LHAASO Observations

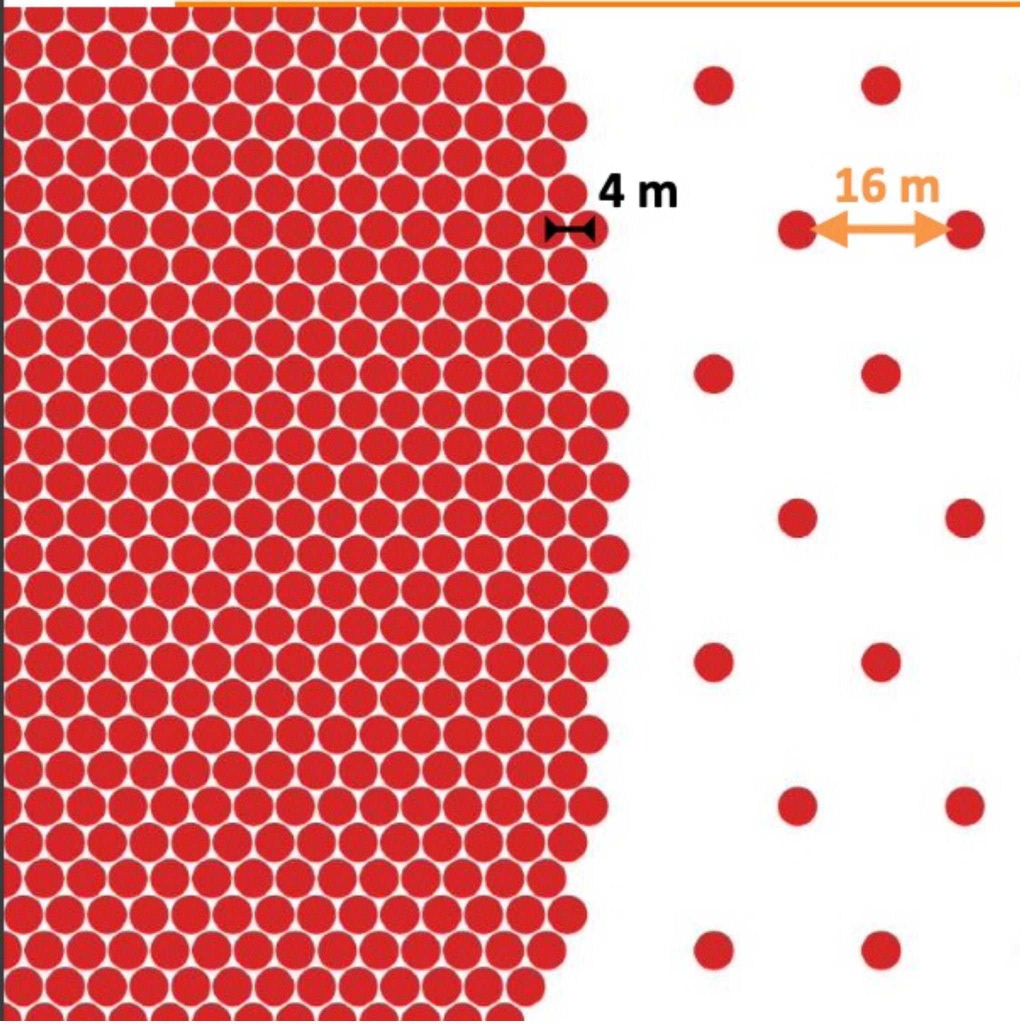


95% one-sided lower limits on DM lifetime obtained with the profile likelihood analysis (thick black lines), for DM decaying into b quarks (left) or τ leptons (right). The black dashed line shows the limit for prompt DM contribution only. The green and yellow bands correspond to the expected 68% and 95% limit ranges from Monte Carlo simulations with the background-only hypothesis. Previous limits are shown with gray and blue lines. The hatched regions show the 1σ DM parameter space favored by IceCube high-energy neutrino flux

Z. Cao et al. (LHAASO Collaboration) Phys. Rev. Lett. 129, 261103 (2022)

SWGO
The Southern Wide-field Gamma-ray Observatory

The baseline detector concept

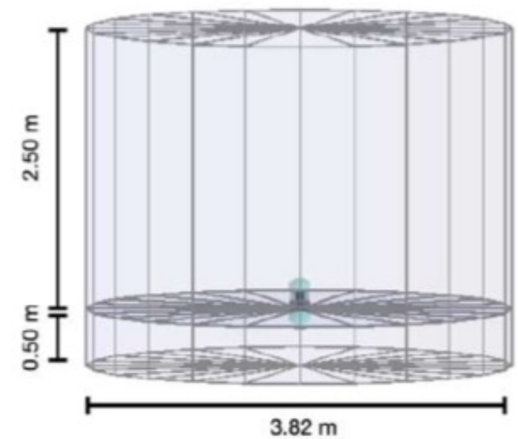


Core: Ø 320 m, FF = 80%
5,700 WCD units

Outer: Ø 600 m, FF = 5%
880 WCD units

Altitude: 4,700 m a.s.l.

✧ muon counting



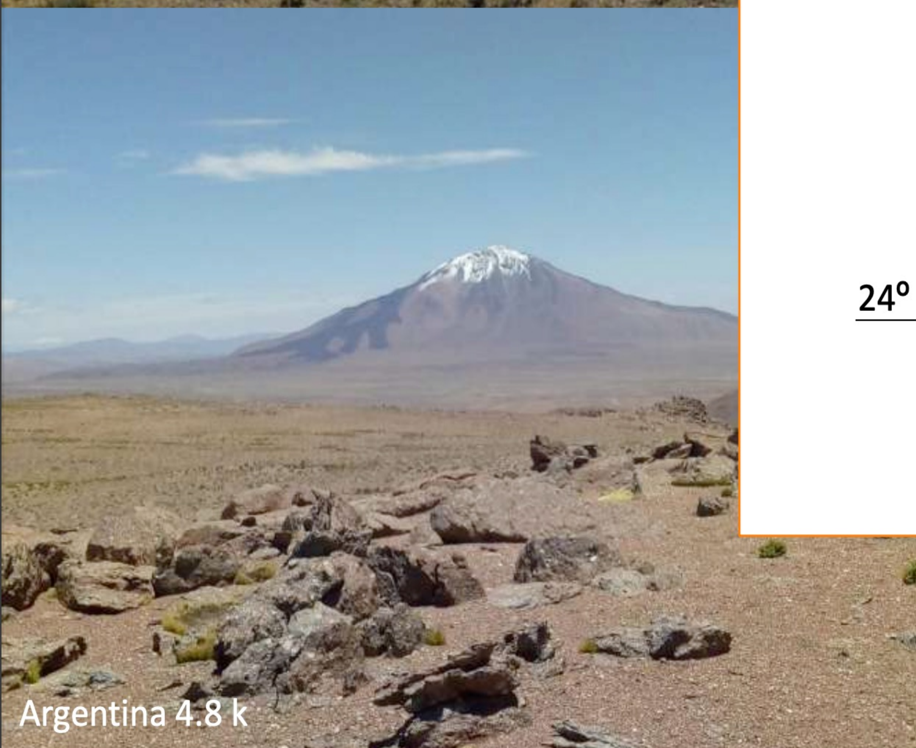
SWGO White paper arXiv:1902.08429

A Wide-field Gamma-ray Observatory in the South

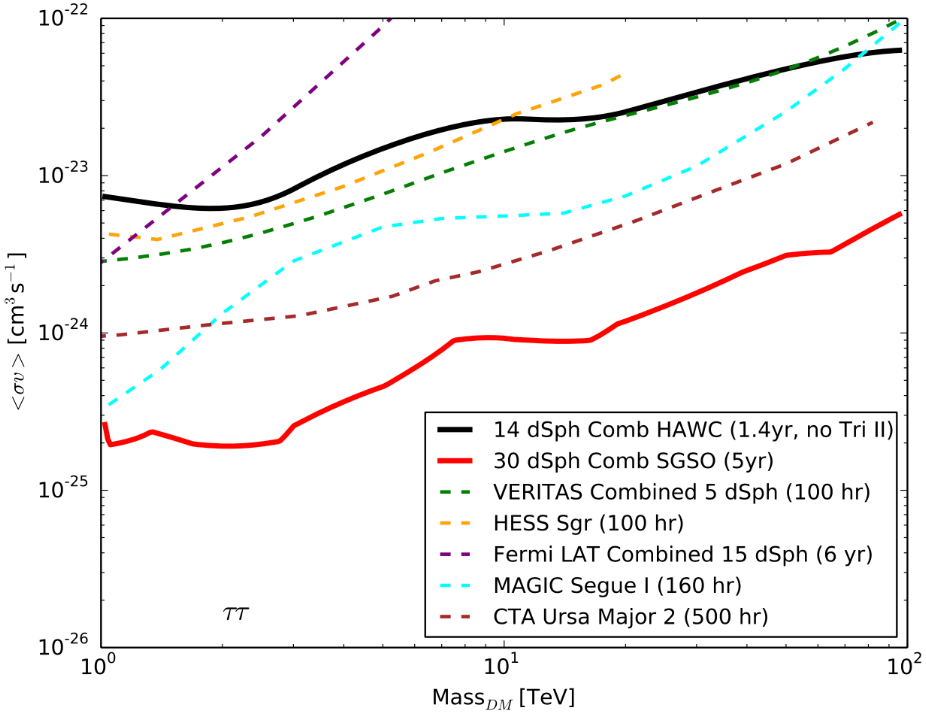
A Wide-field Gamma-ray Observatory in the South

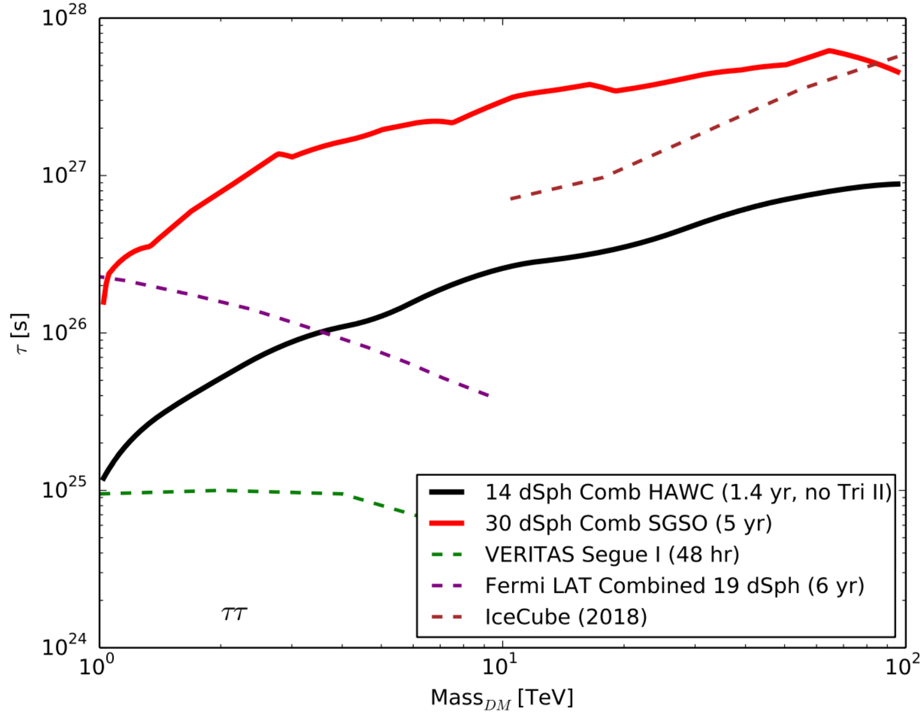
Shortlisting: Fall 2022
Site visits: October
Site selection: Fall 2023

Country	Elevation	Location:
Peru	4900	Laguna Sibinacocha
Peru	4450	Imata
Peru	4450	Yanque
Argentina	4800	Cerro Vecar
Argentina	4450	Alto Tocomar
Chile	4700	ALMA Pampa La Bola
Chile	4400	AAP Pajonales
Bolivia	4700	ALPACA area



SWGO sensitivities





Assumed new dSph discovery and
J-factor and D-factor distributions of the new dSphs matches that of the previously known dSphs

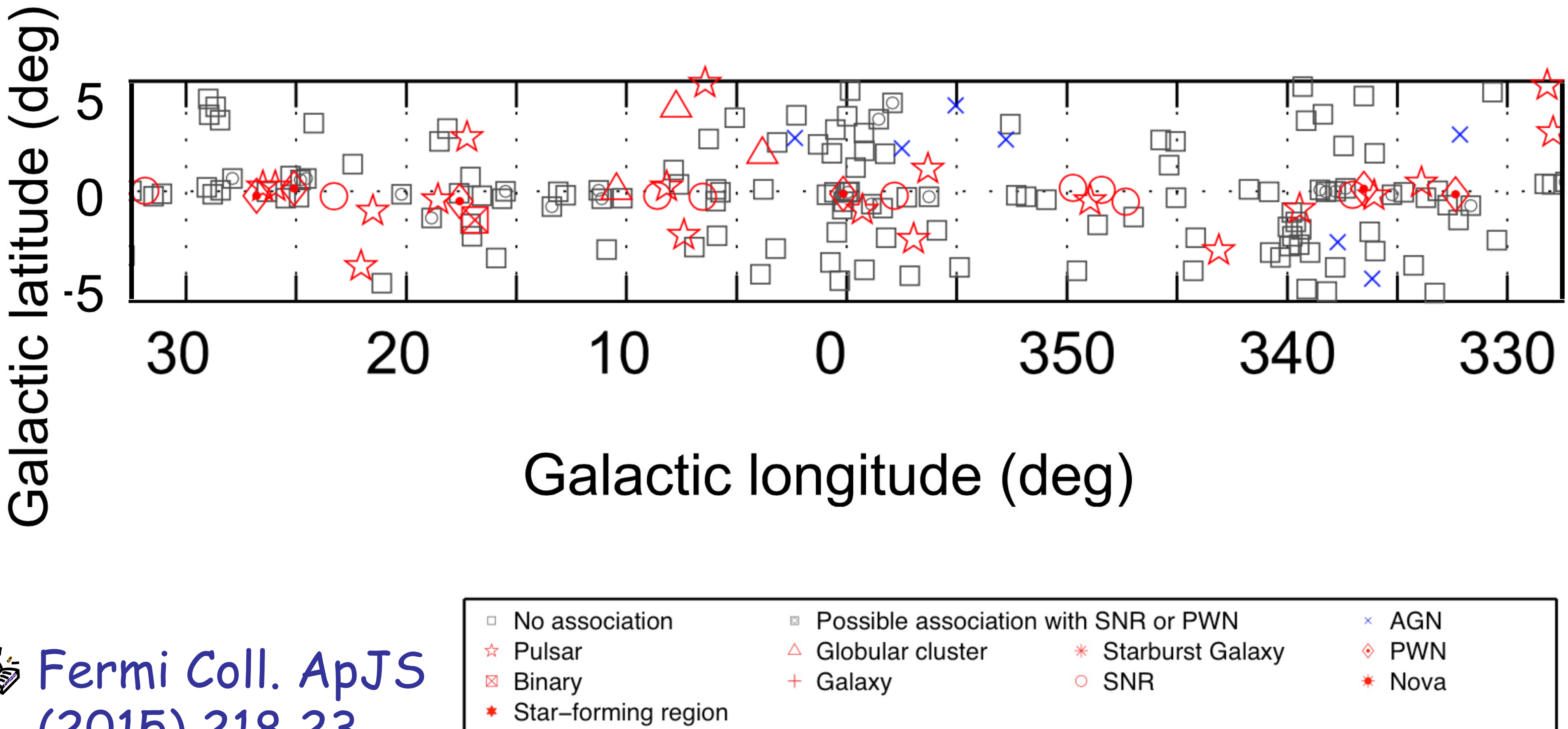
SWGO White paper arXiv:1902.08429

The Low Energy Frontier

The Fermi LAT 3FGL Inner Galactic Region

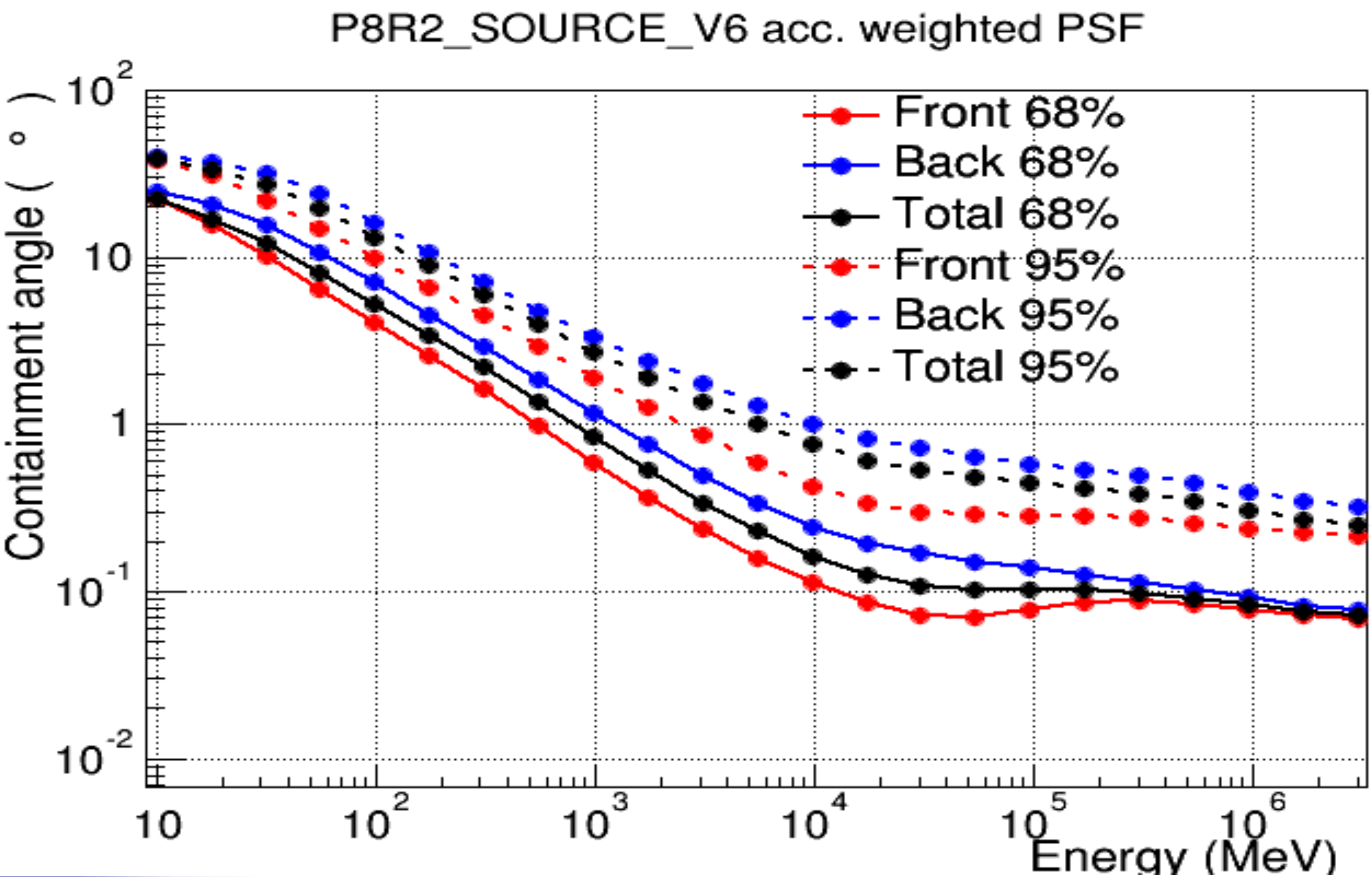
August 4, 2008, to July 31, 2010

100 MeV to 300 GeV energy range

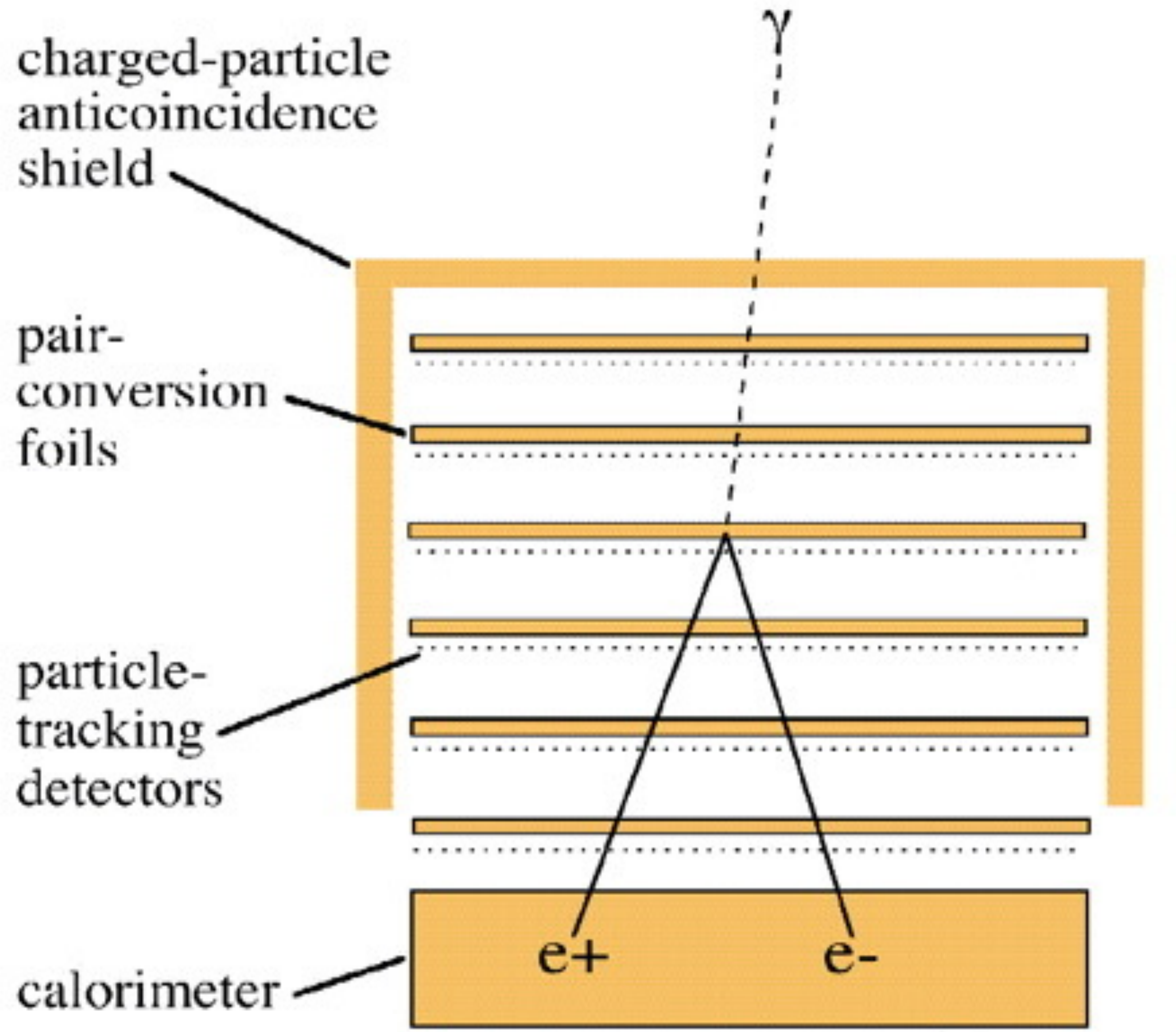


Fermi Coll. ApJS (2015) 218 23
arXiv:1501.02003

Fermi-LAT Instrument Response Functions (Pass 8) Angular Resolution



Elements of a pair-conversion telescope



(energy measurement)

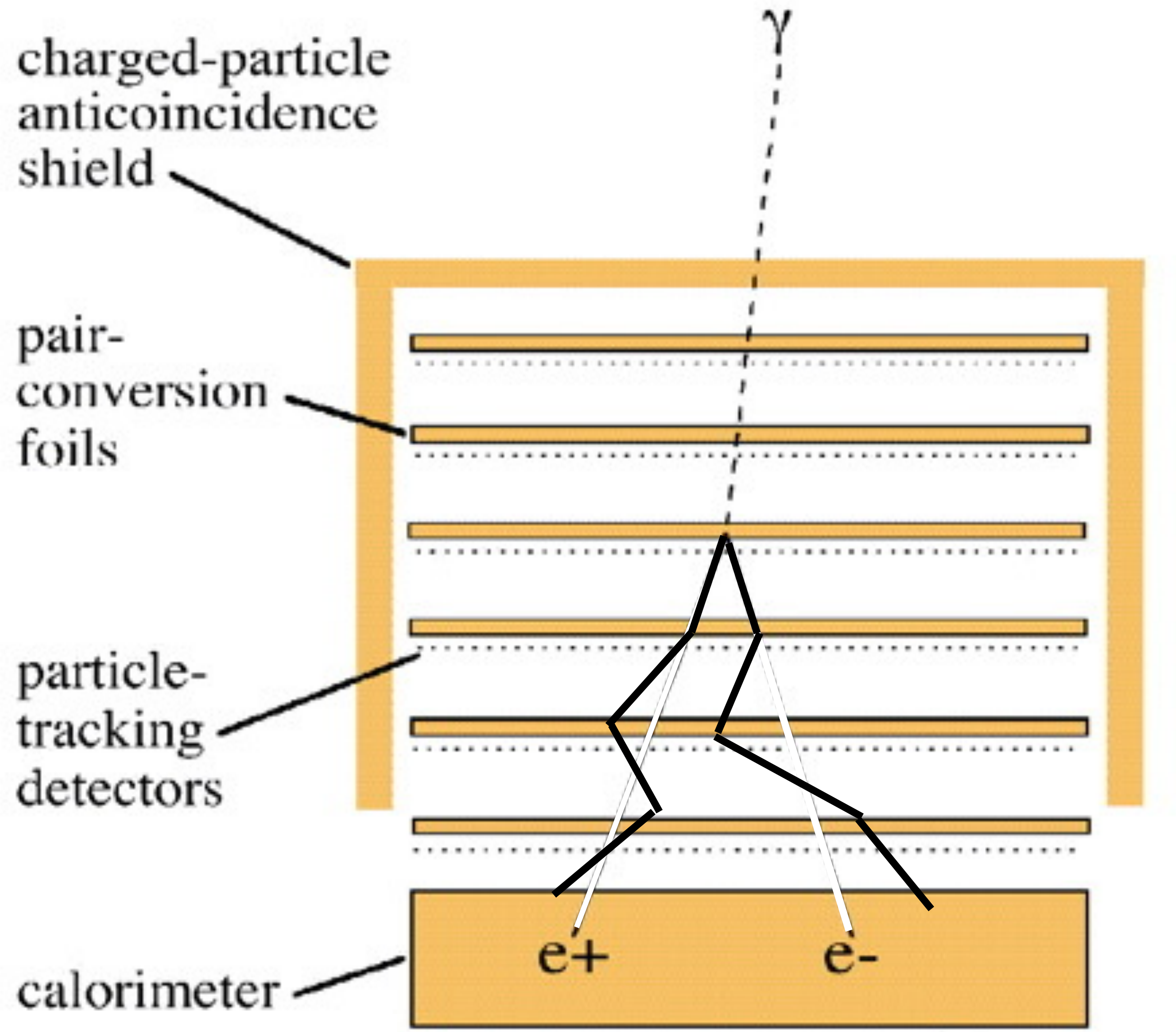
- photons materialize into matter-antimatter pairs:

  E_γ --> $m_{e^+}c^2 + m_{e^-}c^2$

- electron and positron carry information about the direction, energy and polarization of the γ-ray

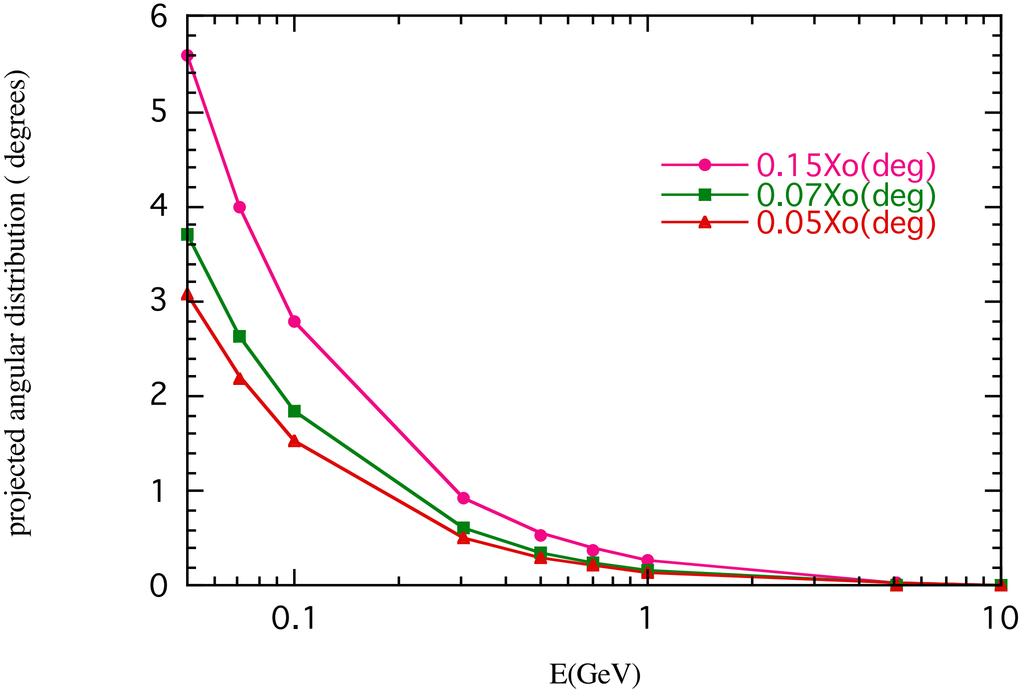
Elements of a pair-conversion telescope

(more realistic scheme)

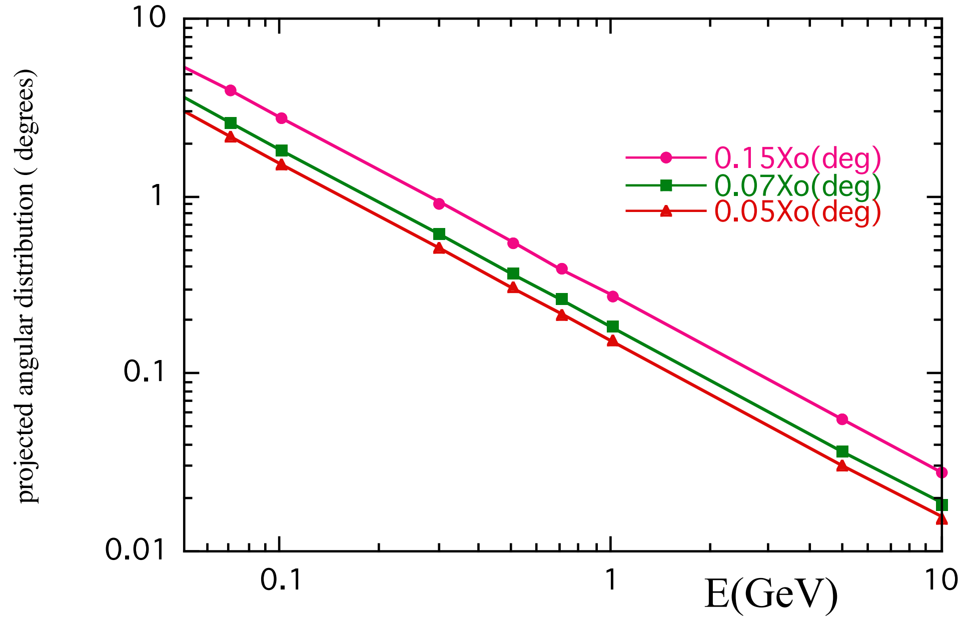


(energy measurement)

- photons materialize into matter-antimatter pairs:
  E_γ --> $m_{e^+}c^2 + m_{e^-}c^2$
- electron and positron carry information about the direction, energy and polarization of the γ-ray



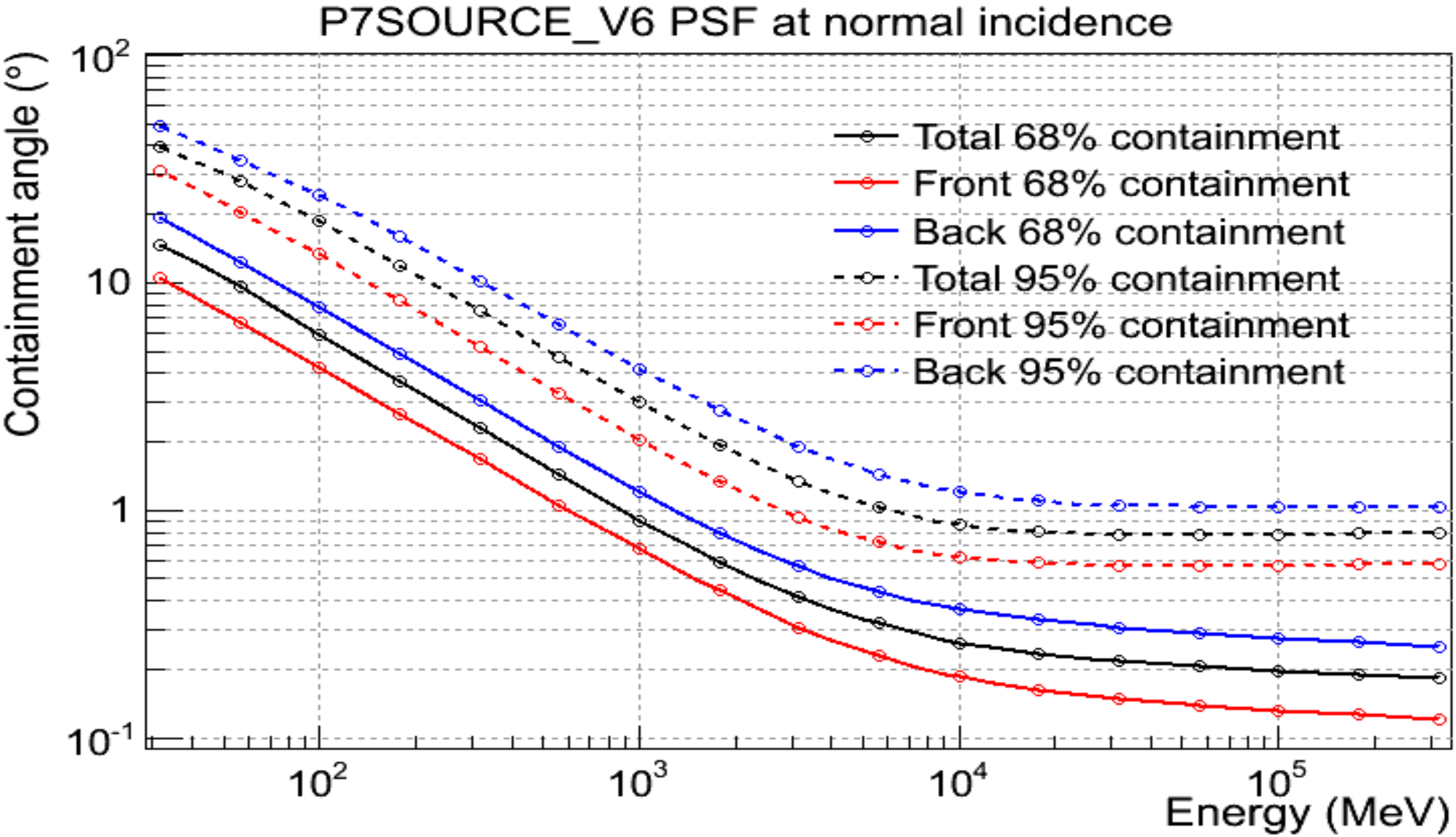
Multiple Scattering



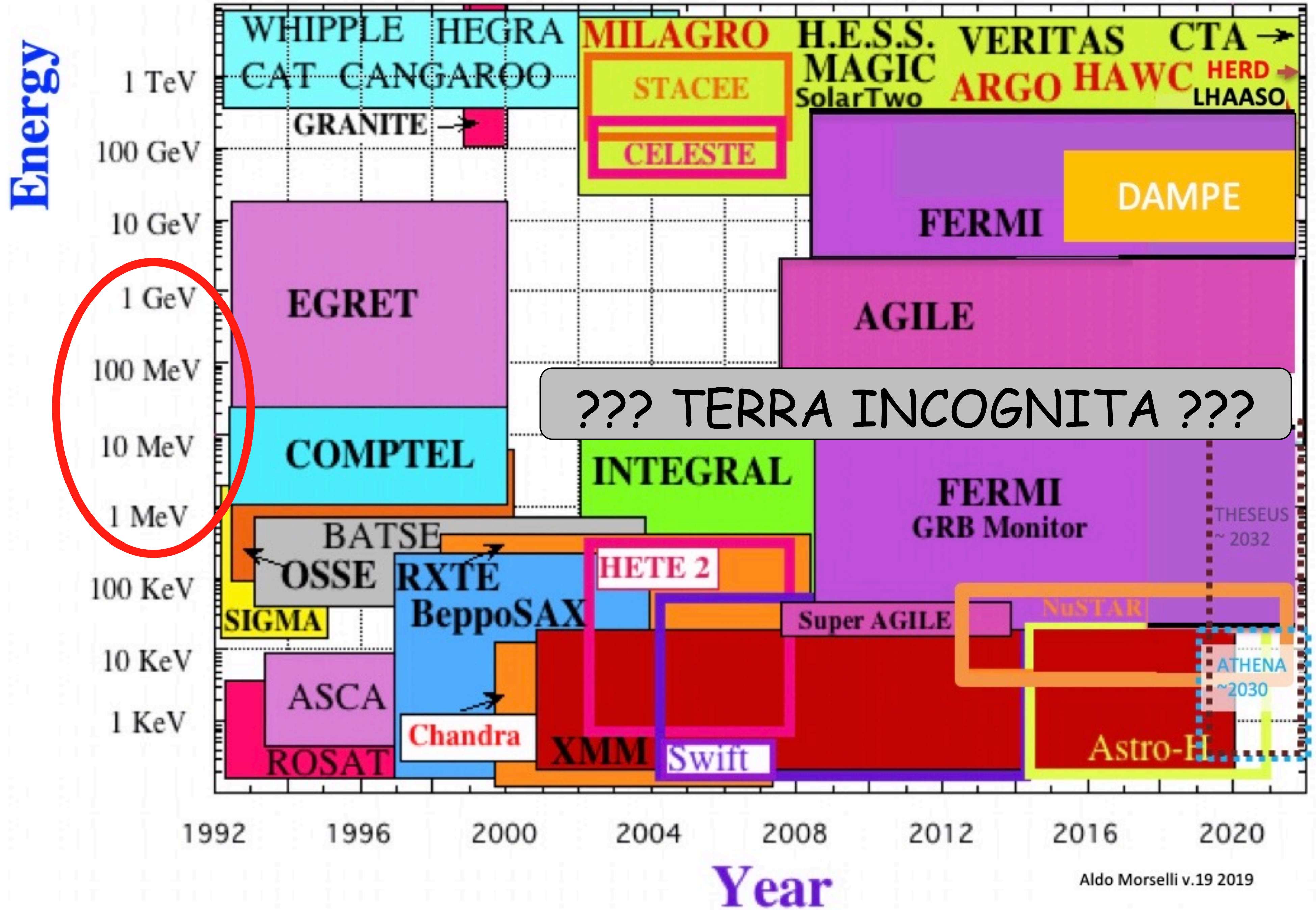
$$\theta_0 = \theta_{plane}^{rms} = \frac{1}{\sqrt{2}}\theta_{space}^{rms}$$

$$\theta_0 = \frac{13.6MeV}{\beta cp} z\sqrt{x/X_0}\left[1 + 0.038\ln\left(x/X_0\right)\right]$$

Fermi Instrument Response Function



http://www.slac.stanford.edu/exp/glast/groups/canda/lat_Performance.htm

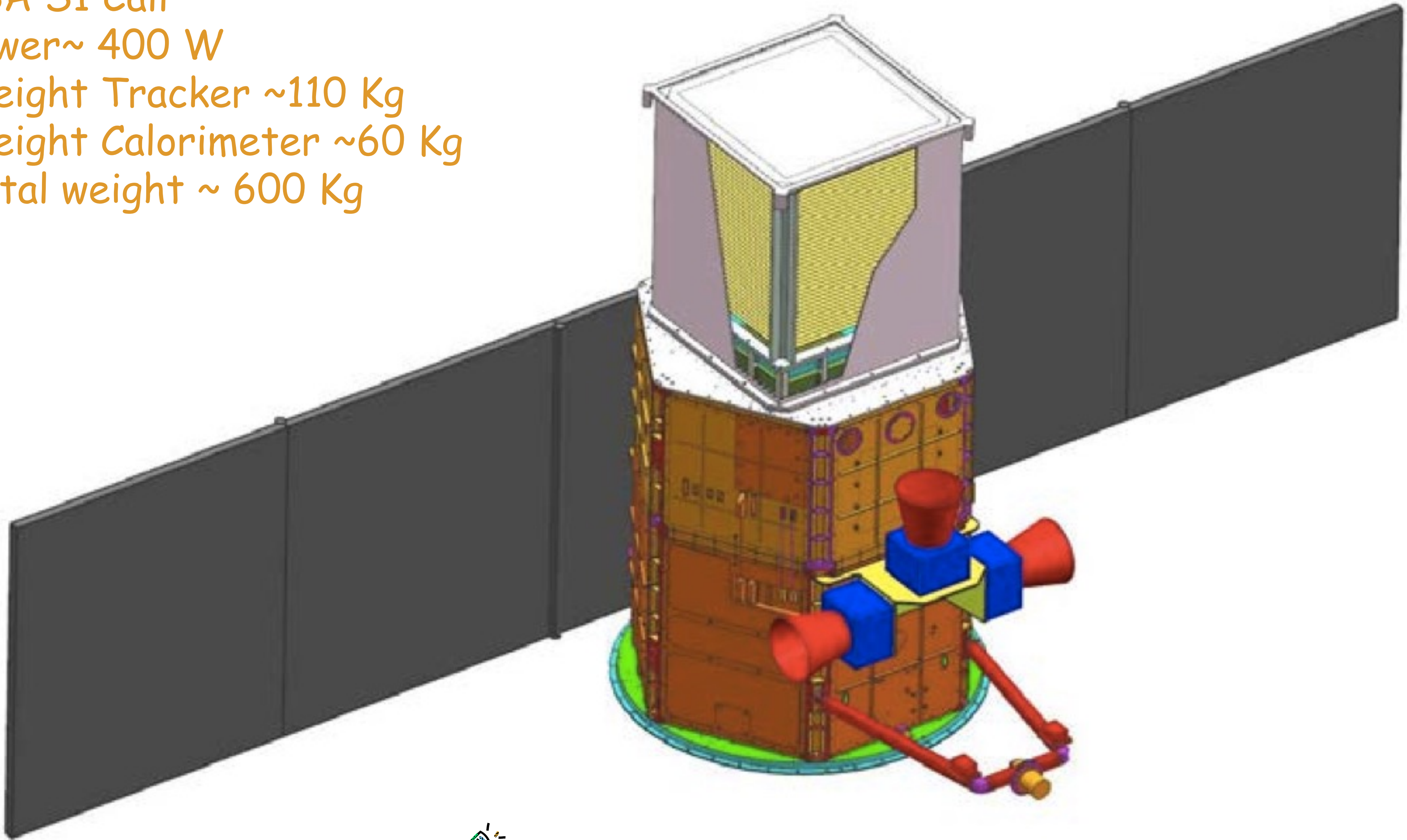


Energy
1 TeV
100 GeV
10 GeV
1 GeV
100 MeV
10 MeV
1 MeV
100 KeV
10 KeV
1 KeV
WHIPPLE
HEGRA
CAT
CANGAROO
GRANITE
MILAGRO
STACEE
CELESTE
H.E.S.S.
MAGIC
SolarTwo
VERITAS
ARGO
HAWC
CTA
HERD
LHAASO
FERMI
DAMPE
EGRET
AGILE
??? TERRA INCOGNITA ???
COMPTEL
INTEGRAL
FERMI
GRB Monitor
THESEUS
~ 2032
BATSE
OSSE
SIGMA
RXTE
BeppoSAX
HETE 2
Super AGILE
NuSTAR
ATHENA
~2030
ASCA
Chandra
XMM
Swift
ROSAT
Astro-H
1992
1996
2000
2004
2008
2012
2016
2020
Year
Aldo Morselli v.19 2019

- 1-100 MeV unexplored domain for
  - Dark Matter searches
  - Galactic compact stars and nucleosynthesis
  - Cosmic rays
  - Relativistic jets, microquasars
  - Blazars
  - Gamma-Ray Bursts
  - Solar physics
- and…
  - Terrestrial Gamma-Ray Flashes

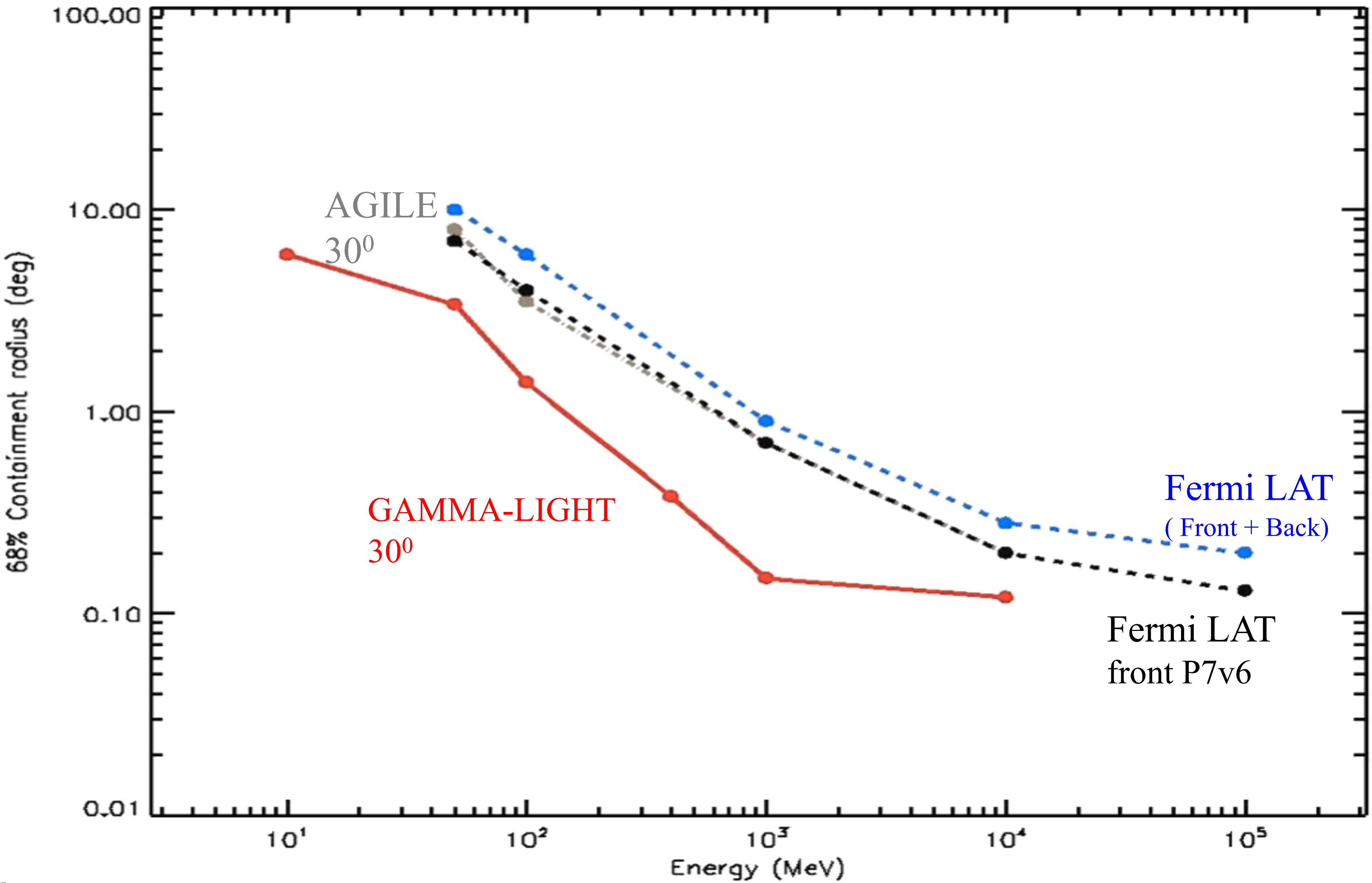
Gamma-light project

ESA S1 Call
Power~ 400 W
Weight Tracker ~110 Kg
Weight Calorimeter ~60 Kg
Total weight ~ 600 Kg



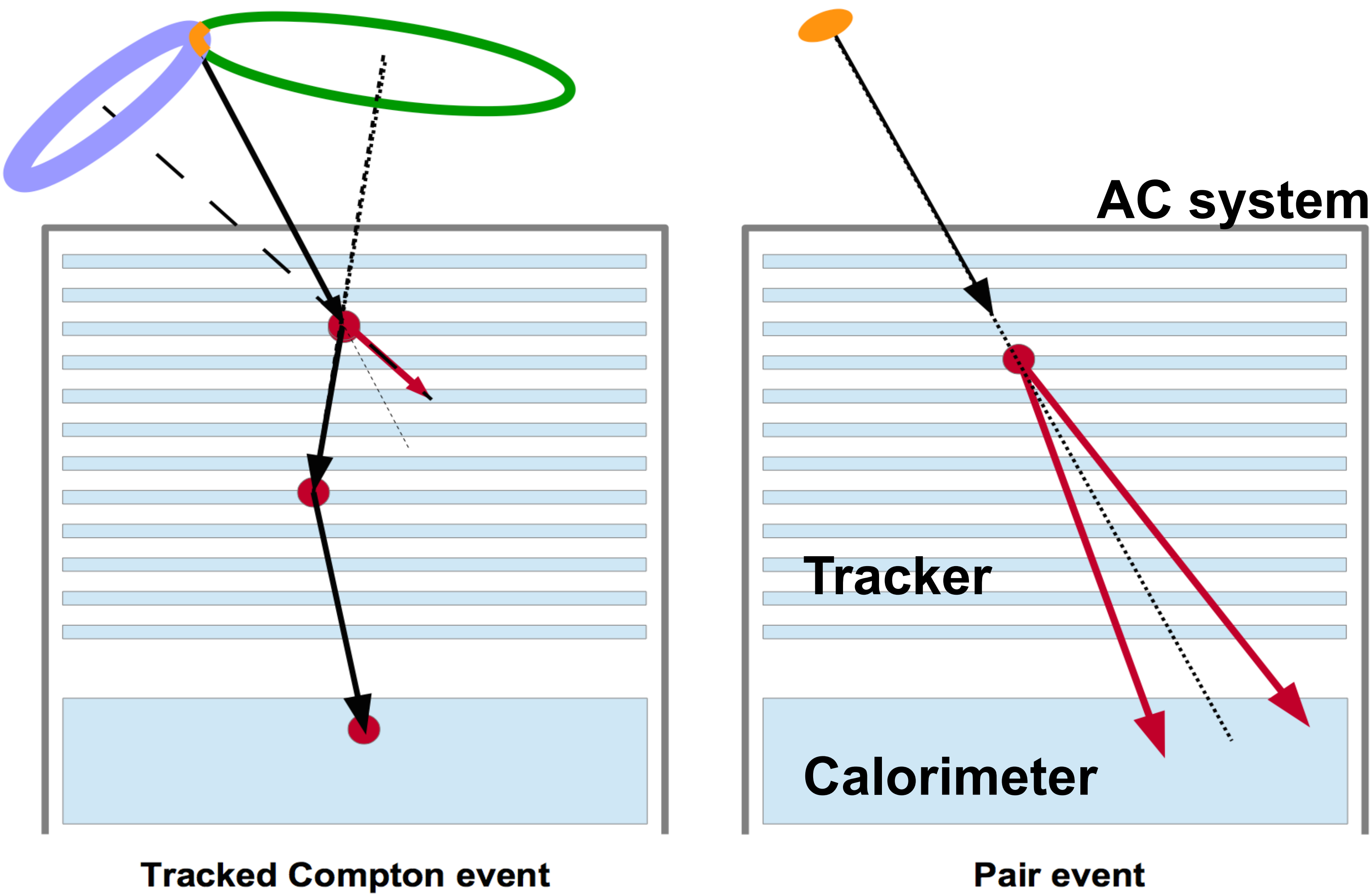
A.Morselli et al. , Nuclear Physics B Proc. Supp. 239–240 (2013) 193-198 [arXiv:1406.1071]

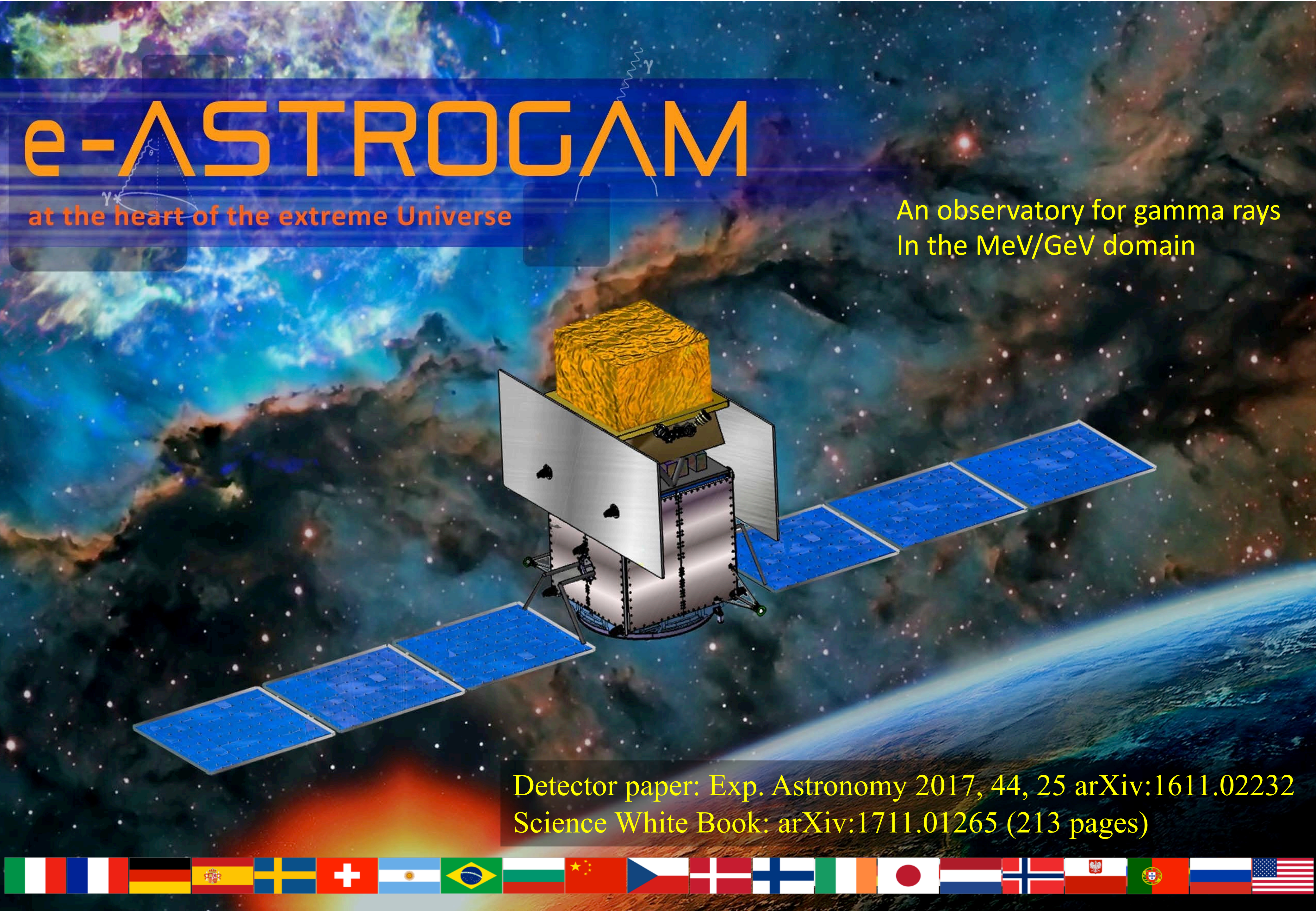
Gamma-Light Point Spread Function (angular resolution)



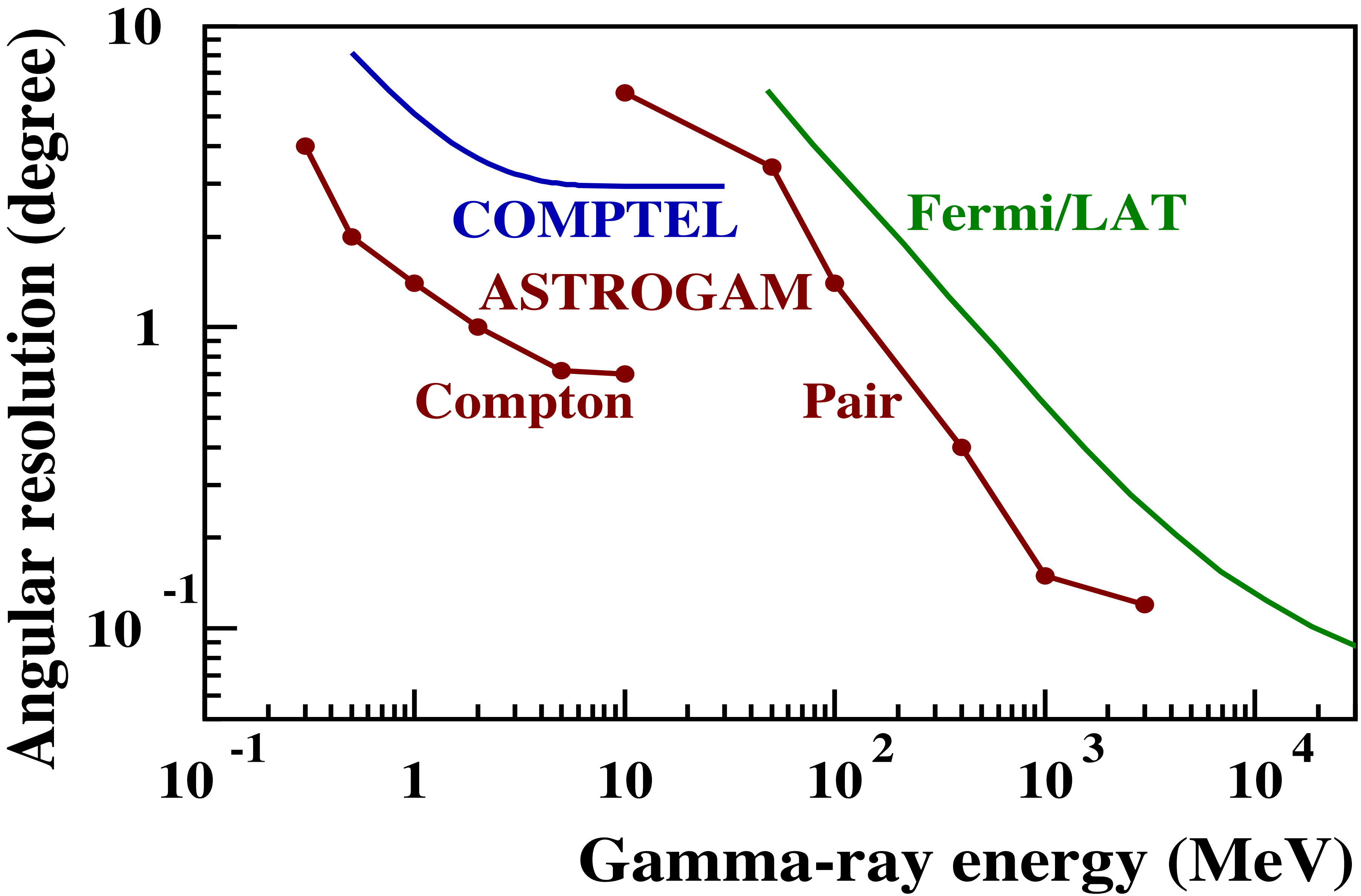
A.Morselli et al. , Nuclear Physics B Proc. Supp. 239–240 (2013) 193-198 [arXiv:1406.1071]

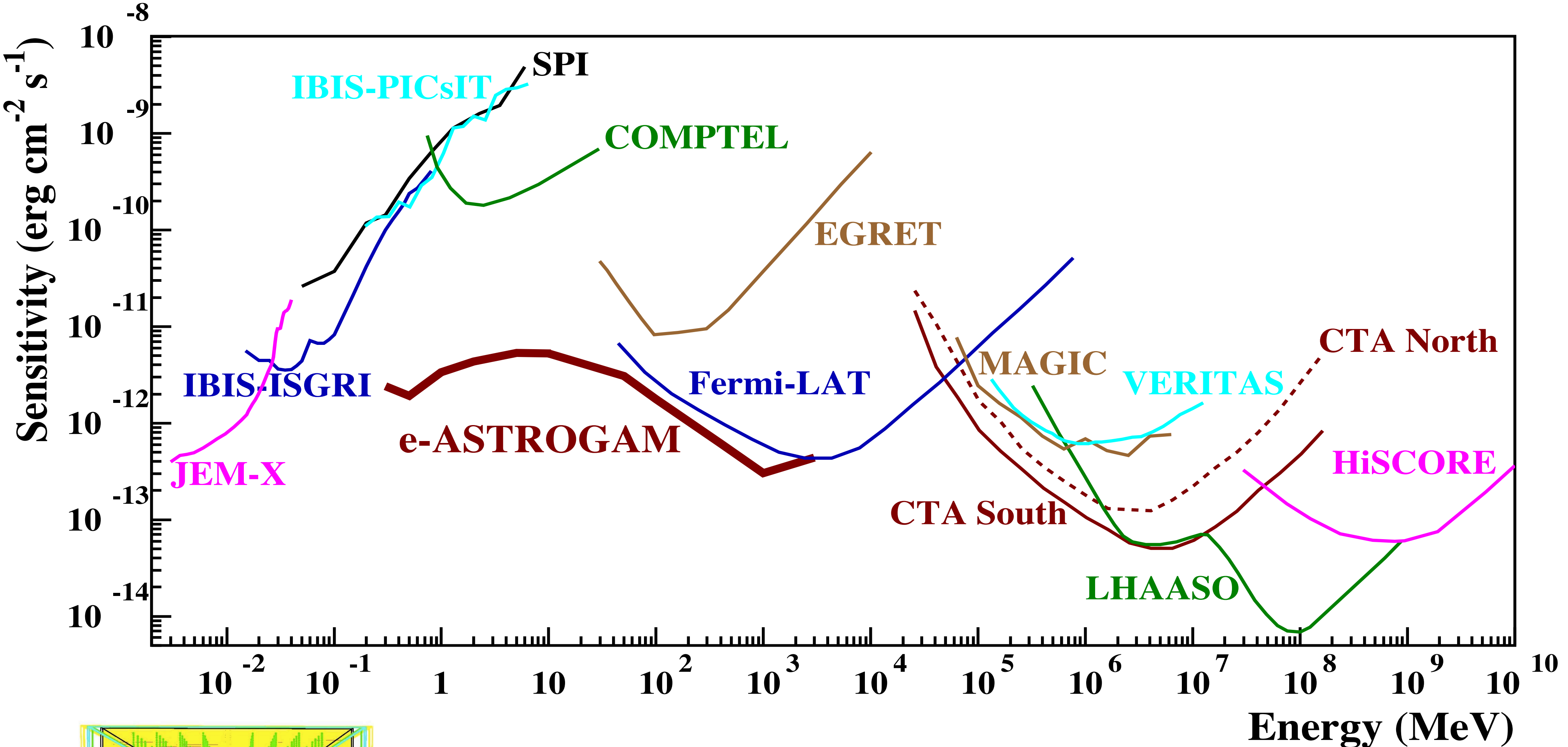
An instrument that combine two detection techniques

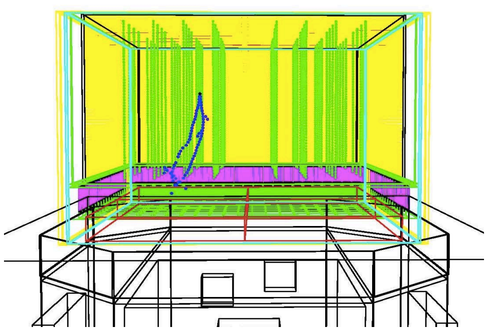



e-ASTROGAM
at the heart of the extreme Universe
An observatory for gamma rays
In the MeV/GeV domain
Detector paper: Exp. Astronomy 2017, 44, 25 arXiv:1611.02232
Science White Book: arXiv:1711.01265 (213 pages)

ASTROGAM Angular Resolution





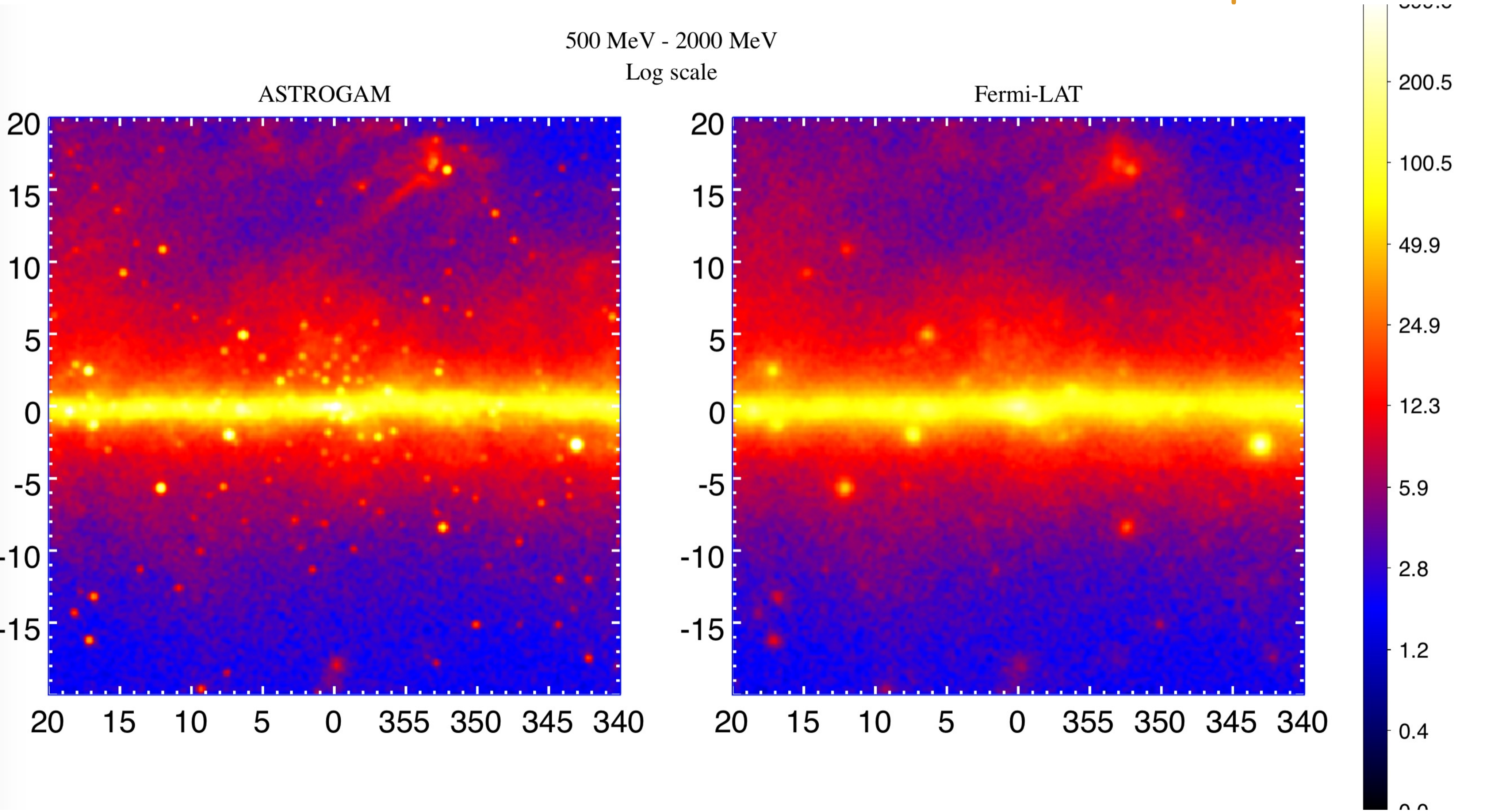


- e-ASTROGAM performance evaluated with **MEGAlib** and – both tools based on Geant4 – and a detailed numerical mass model of the gamma-ray instrument

e-Astrogam: arXiv:1611.02232

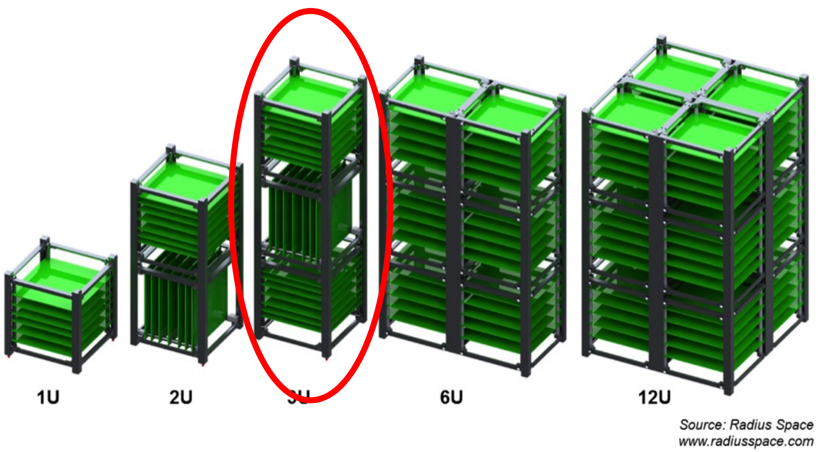
Galactic Center Region 0.5-2 GeV

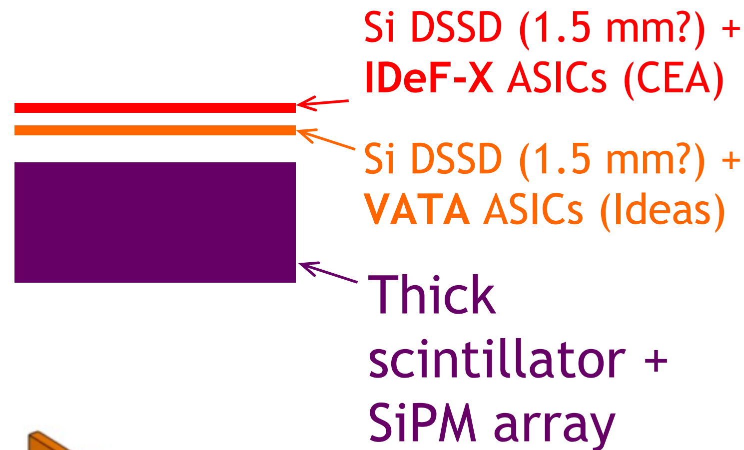
Fermi PSF Pass7 rep v15 source



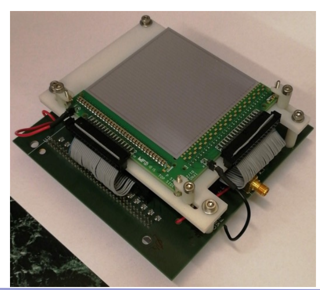
COMCUBE Nanosat sub-WP

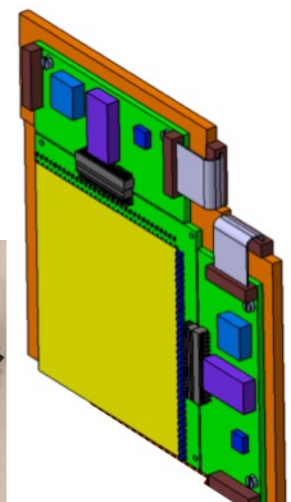
Development of a 3U (?) Compton nanosat for the polarimetry of GRBs + qualification of the e-ASTROGAM technologies

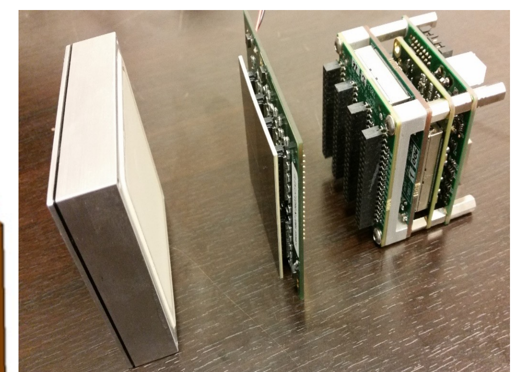




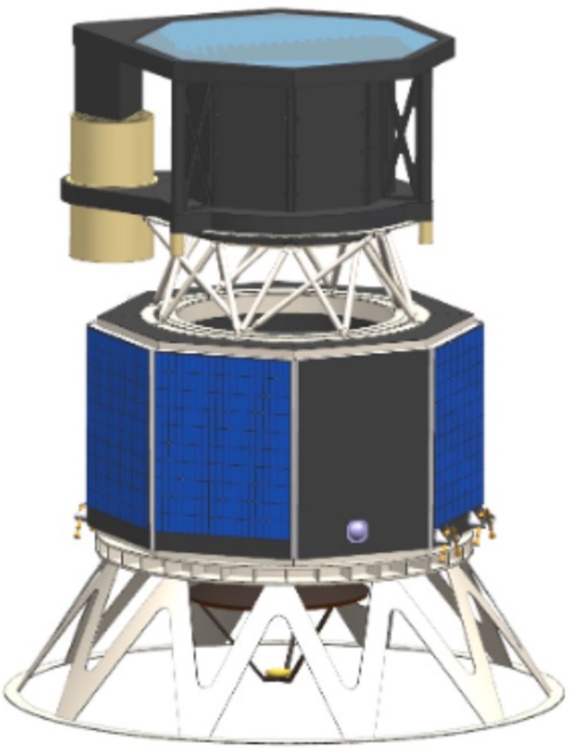
- Cubesat : standard unit ⇒ 1U
- Size : 10 x 10 x 10 cm
- Weight : 1kg
- Power: ~ 1.3 W

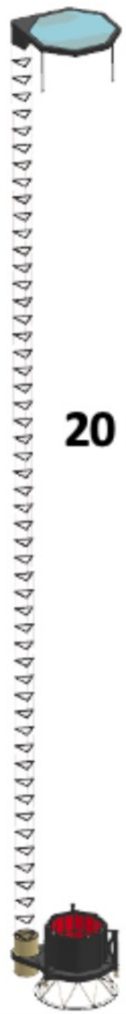


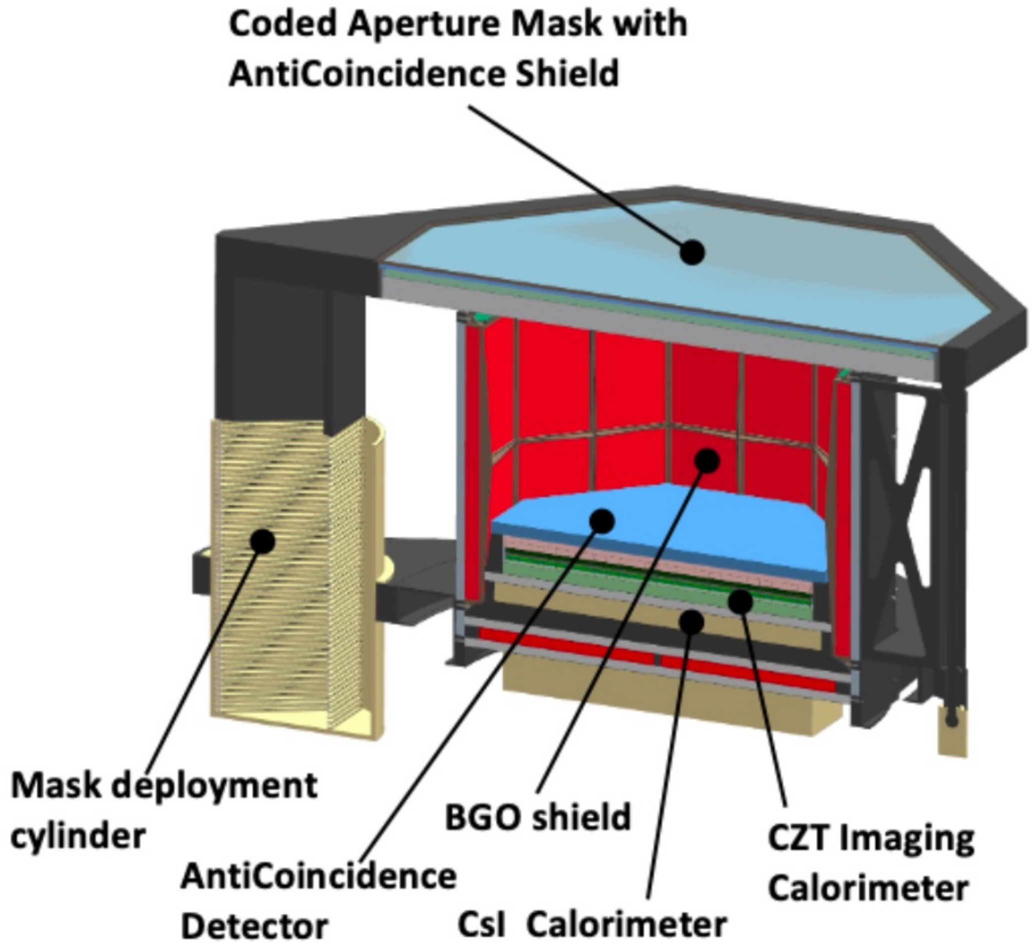




GECCO The Galactic Explorer with a Coded Aperture Mask Compton Telescope



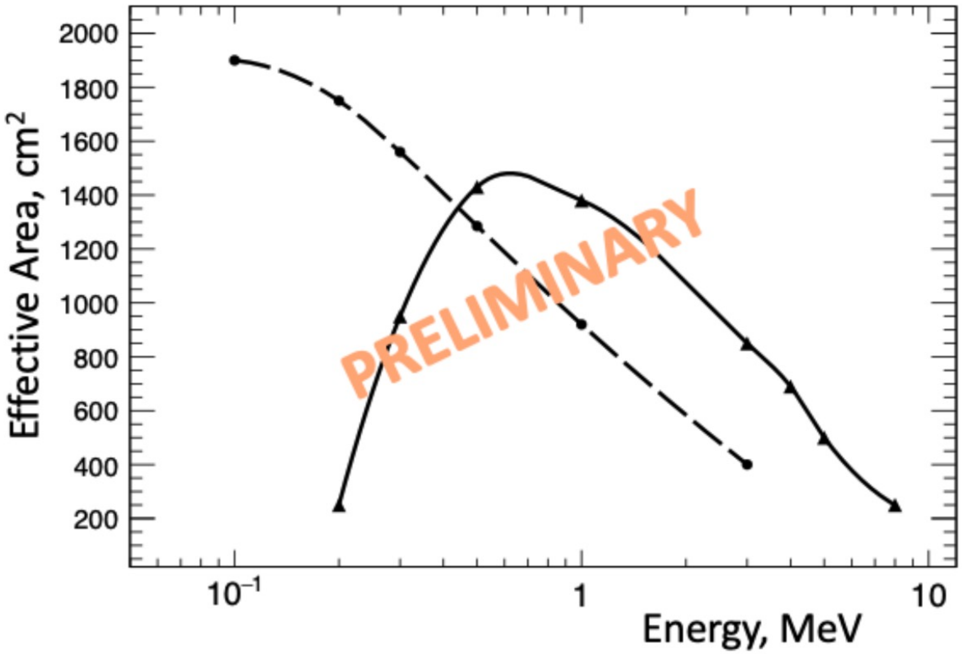




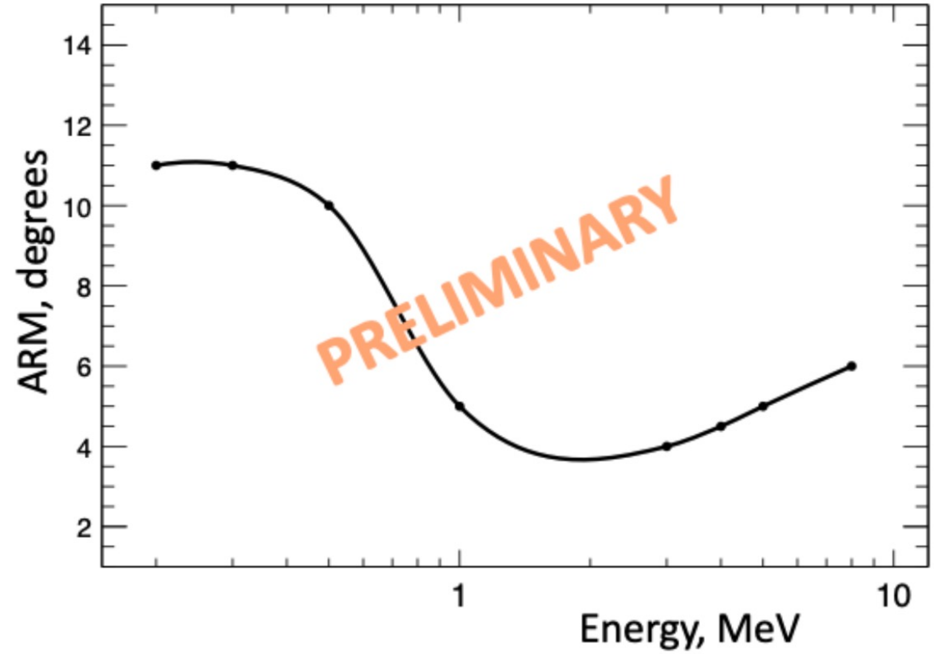
conceptual design | mask in stowed position | cutaway | diameter =90 cm

GECCO Team, JCAP07(2022)036 arXiv:2112.07190

GECCO The Galactic Explorer with a Coded Aperture Mask Compton Telescope



effective area for the CA mask imaging; the solid line is for Compton pointing used, and the dashed line is for classical mask analysis.

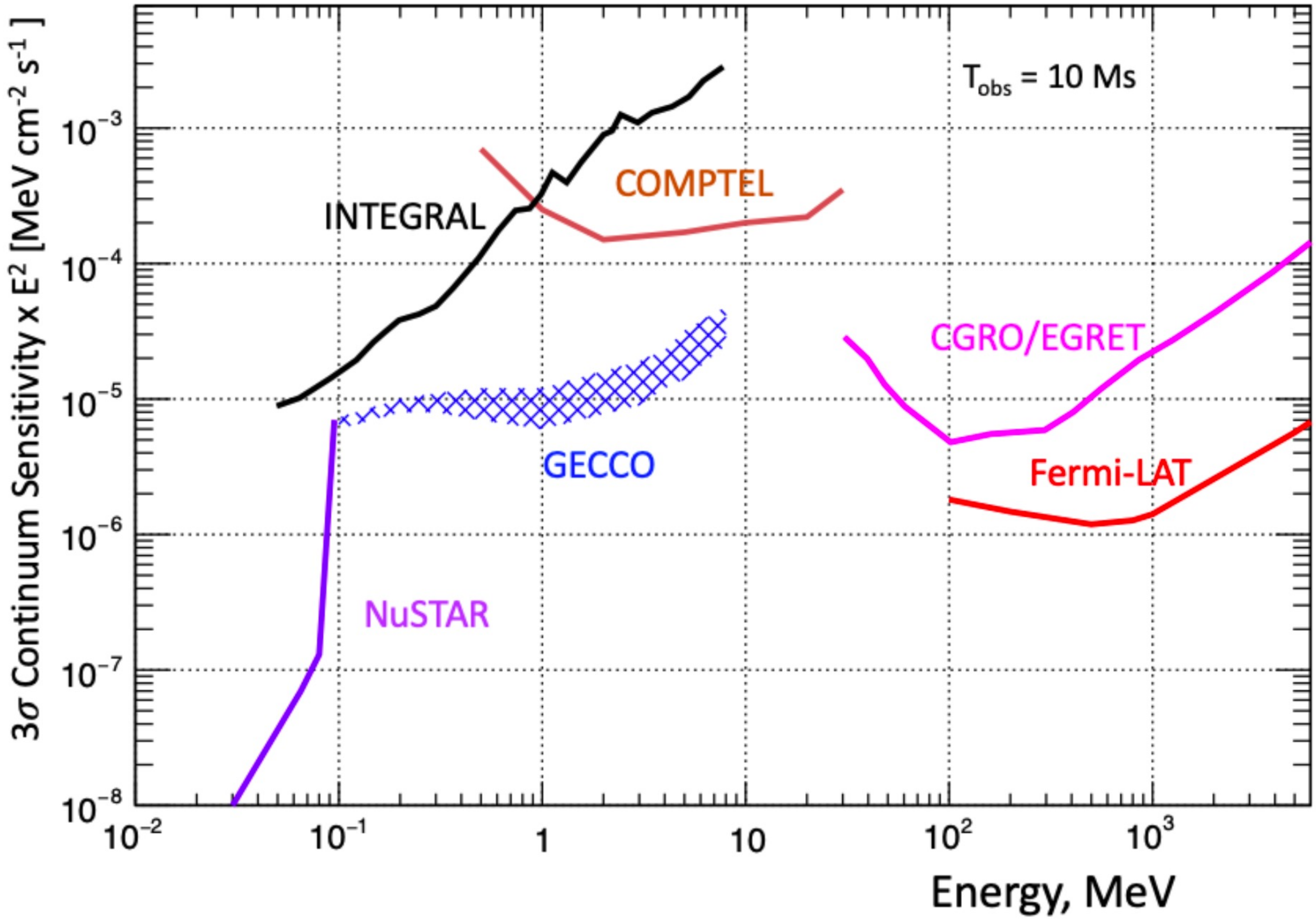


ARM (angular resolution measure) for the ImCal standalone Compton telescope.

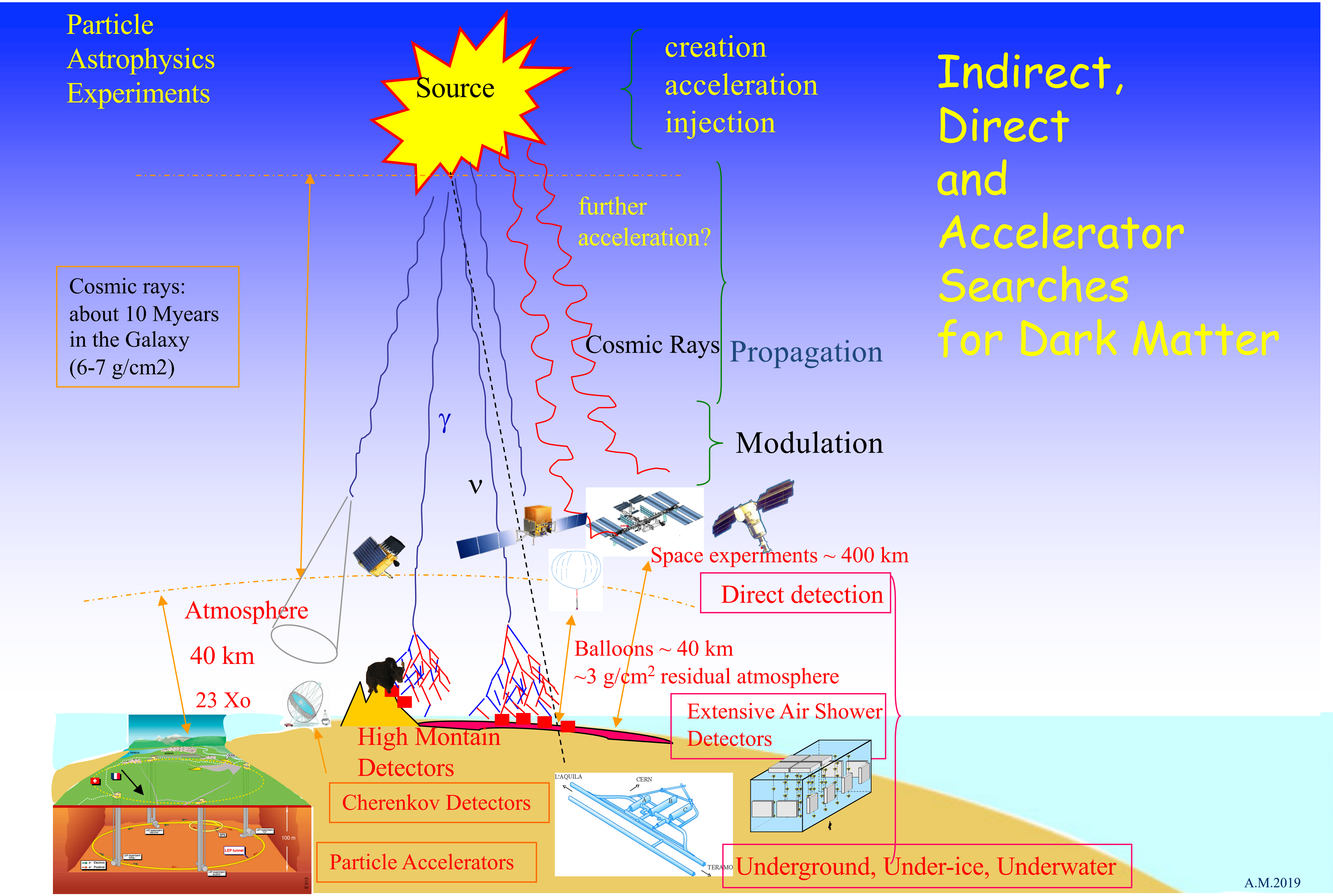
GECCO Team, in preparation

GECCO The Galactic Explorer with a Coded Aperture Mask Compton Telescope

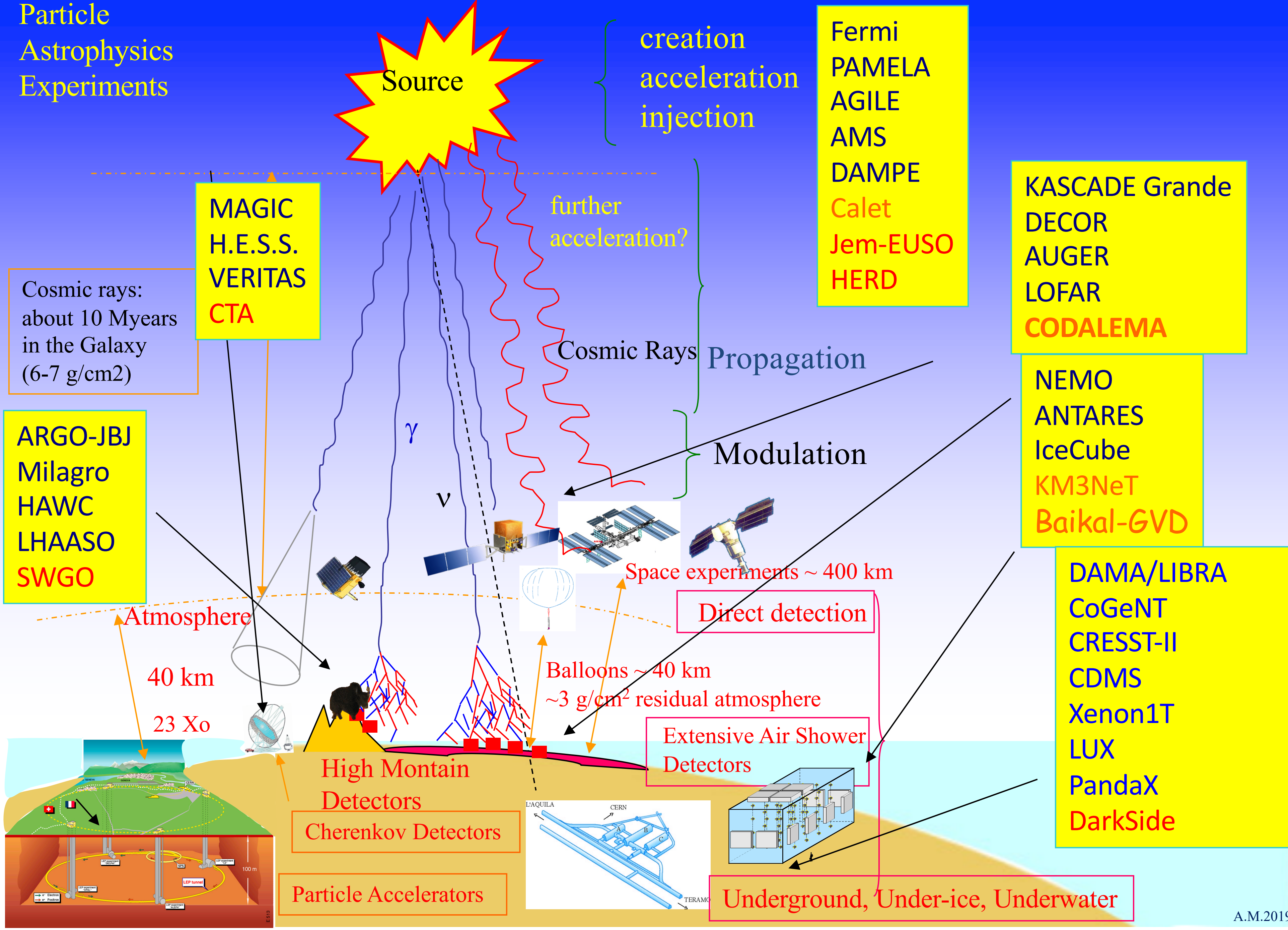
Sensitivity



GECCO Team, JCAP07(2022)036 arXiv:2112.07190



Particle Astrophysics Experiments
Source
creation
acceleration
injection
Indirect, Direct and Accelerator Searches for Dark Matter
further acceleration?
Cosmic rays: about 10 Myears in the Galaxy (6-7 g/cm2)
Cosmic Rays
Propagation
γ
Modulation
ν
Space experiments ~ 400 km
Direct detection
Atmosphere
40 km
23 Xo
Balloons ~ 40 km
~3 g/cm^2 residual atmosphere
Extensive Air Shower Detectors
High Montain Detectors
Cherenkov Detectors
Particle Accelerators
L'AQUILA
CERN
TERAMO
Underground, Under-ice, Underwater
A.M.2019


Particle
Astrophysics
Experiments
Source
creation
acceleration
injection
Fermi
PAMELA
AGILE
AMS
DAMPE
Calet
Jem-EUSO
HERD
MAGIC
H.E.S.S.
VERITAS
CTA
further
acceleration?
KASCADE Grande
DECOR
AUGER
LOFAR
CODALEMA
Cosmic rays:
about 10 Myears
in the Galaxy
(6-7 g/cm2)
Cosmic Rays
Propagation
NEMO
ANTARES
IceCube
KM3NeT
Baikal-GVD
ARGO-JBJ
Milagro
HAWC
LHAASO
SWGO
γ
Modulation
ν
Space experiments ~ 400 km
DAMA/LIBRA
CoGeNT
CRESST-II
CDMS
Xenon1T
LUX
PandaX
DarkSide
Atmosphere
Direct detection
40 km
Balloons ~ 40 km
~3 g/cm² residual atmosphere
23 Xo
Extensive Air Shower
Detectors
High Montain
Detectors
Cherenkov Detectors
L'AQUILA
CERN
TERAMO
Particle Accelerators
Underground, Under-ice, Underwater
A.M.2019