Cosmological gravitational particle production of massive spin-2 particles

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based on arXiv:2302.04390, accepted by JHEP

Table of contents

- Gravitational particle production
- Massive spin-2 field in Hassan-Rosen bigravity
- Results:
  - Number density & relic abundance of produced particles
  - Constraints on parameter space: mass & reheating temperature

Problem

- Massive spin-2 particles is a DM candidate: arXiv:1607.03497.
- Previously, no comprehensive work on how it can be produced gravitationally.
- How does it differ from other CGPP scenarios?

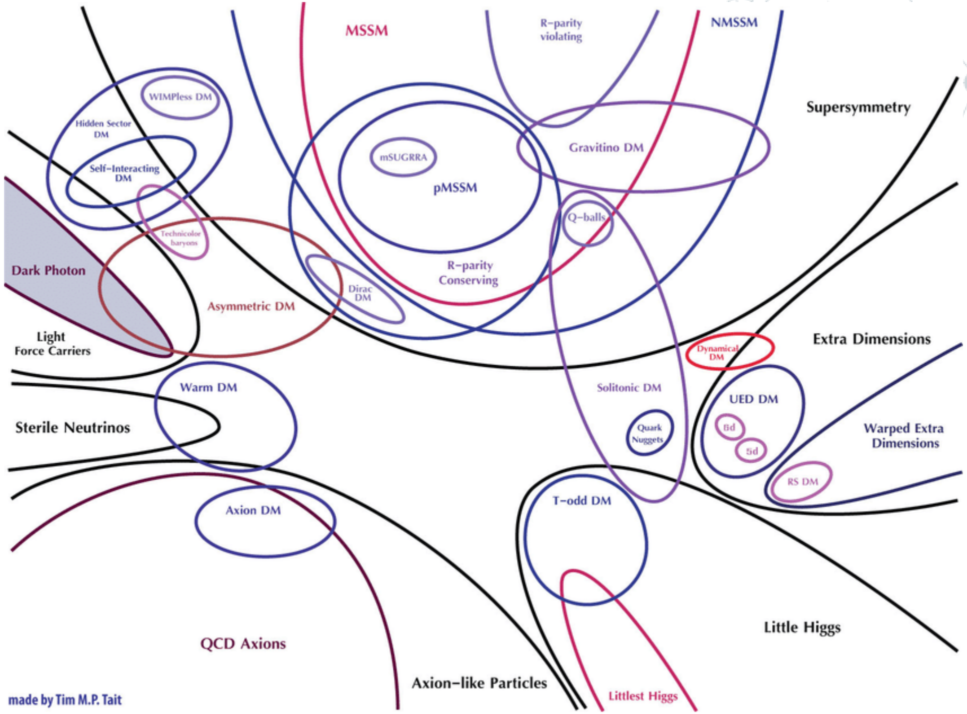


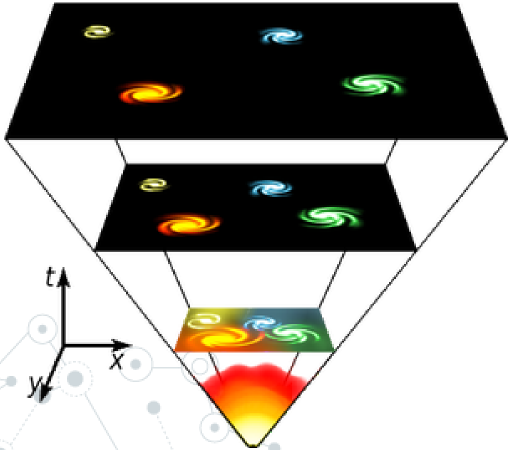
Figure due to Timothy Tait

Gravitational particle production

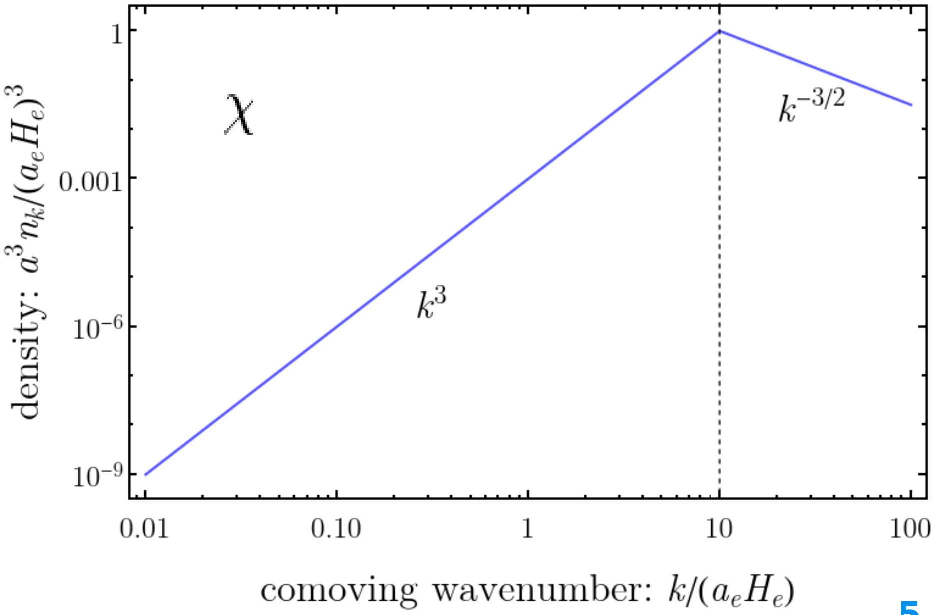
Gravitational particle production: big picture

- Inflationary gravitational background can amplify vacuum fluctuations of matter fields.
- Produce a spectrum of produced particles peaked at Hubble scale of inflation.

$$S = \int d^4x \sqrt{-g}\, \mathcal{L}_{\text{dark matter}}(g_{\mu\nu}, \chi)$$



induces



Gravitational particle production: example

Cosmological expansion → Time-dependent Lagrangian → Time-dependent effective frequency → Particle production

$$ds^2 = a(\eta)^2(-\mathrm{d}\eta^2 + \mathrm{d}\mathbf{x}^2) \quad \Rightarrow \quad \mathcal{L} = a^2\left[\frac{1}{2}(\partial_\eta\phi)^2 - \frac{1}{2}(\nabla\phi)^2 - \frac{1}{2}a^2m^2\phi^2\right]$$

$$\Downarrow$$

$$|\beta_k|^2 = \frac{\omega_k\left|a\tilde{\phi}_k\right|^2}{2} + \frac{\left|\partial_\eta(a\tilde{\phi}_k)\right|^2}{2\omega_k} - \frac{1}{2} \quad \Leftarrow \quad [\partial_\eta^2 - \nabla^2 + a^2(\eta)m_{\text{eff}}^2(\eta)](a\phi) = 0$$

$$n_k(\eta) = a(\eta)^{-3}\frac{k^3}{2\pi^2}|\beta_k|^2$$

particle production!

Birrell & Davies (1982), Parker & Toms (2009), L H Ford (2021), ...

Massive spin-2 particles from bigravity

Hassan Rosen Bigravity

Two interacting metrics (spin-2 dofs)

$$S = \int d^4x \Bigg[\underbrace{\frac{M_g^2}{2}\sqrt{-g}\,R[g] + \frac{M_f^2}{2}\sqrt{-f}\,R[f]}_{\text{Einstein-Hilbert}} \underbrace{-m^2 M_*^2 \sqrt{-g}\,V(\mathbb{X};\beta_n)}_{\text{interaction between metrics}}$$

$$\underbrace{+\sqrt{-g}\,\mathcal{L}_g(g,\phi_g) + \sqrt{-f}\,\mathcal{L}_f(f,\phi_f)}_{\text{minimal theory}} \underbrace{+\sqrt{-g_\star}\,\mathcal{L}_\star(g_\star,\phi_\star)}_{\text{nonminimal theory}} \Bigg] .$$

Matter sector: the inflatons

Bigravity: mass eigen modes

- Expand two metrics around same background
- Mass eigen modes can be identified: one massless, one massive

$$g_{\mu\nu} = \bar{g}_{\mu\nu} + \frac{2}{M_g} h_{\mu\nu}\,, \qquad f_{\mu\nu} = \bar{g}_{\mu\nu} + \frac{2}{M_f} k_{\mu\nu}$$

$$\frac{u_{\mu\nu}}{M_*} = \frac{h_{\mu\nu}}{M_f} + \frac{k_{\mu\nu}}{M_g}\,, \qquad \frac{v_{\mu\nu}}{M_*} = \frac{h_{\mu\nu}}{M_g} - \frac{k_{\mu\nu}}{M_f}$$

Massless spin-2 (GW)

Massive spin-2

Bigravity: spin-2 dofs

- Massless spin-2: 2 dofs (2 tensor modes, +/x)
- Massive spin-2: 2s+1=5 dofs (2 tensor & 2 vector & 1 scalar modes)
- Massless / massive modes decouple at quadratic order

$$S = \int \mathrm{d}^4x \left[\sqrt{-\bar{g}}\, \bar{\mathcal{L}}(\bar{g}, \bar{\phi}) + \sqrt{-\bar{g}}\, \mathcal{L}^{(2)}_{\mathrm{massless}} + \sqrt{-\bar{g}}\, \mathcal{L}^{(2)}_{\mathrm{massive}} + \mathrm{interactions} \right]$$

$$\frac{u_{\mu\nu}}{M_*} = \frac{h_{\mu\nu}}{M_f} + \frac{k_{\mu\nu}}{M_g}, \qquad \frac{v_{\mu\nu}}{M_*} = \frac{h_{\mu\nu}}{M_g} - \frac{k_{\mu\nu}}{M_f}$$

The “Minimal” theory

- Each inflaton couples to one metric.
- The “massive” and “massless” sector decouples at quadratic order.

	Original field variables	Mass eigenstates
Metric	$g_{\mu\nu} = \bar{g}_{\mu\nu} + \frac{2}{M_g} h_{\mu\nu}\,, \quad f_{\mu\nu} = \bar{g}_{\mu\nu} + \frac{2}{M_f} k_{\mu\nu}$	$\frac{u_{\mu\nu}}{M_*} = \frac{h_{\mu\nu}}{M_f} + \frac{k_{\mu\nu}}{M_g}\,, \quad \frac{v_{\mu\nu}}{M_*} = \frac{h_{\mu\nu}}{M_g} - \frac{k_{\mu\nu}}{M_f}$
Inflaton	$\phi_g = \bar{\phi}_g + \varphi_g\,, \quad \phi_f = \bar{\phi}_f + \varphi_f$	$\frac{\varphi_u}{M_*} = \frac{\varphi_g}{M_f} + \frac{\varphi_f}{M_g}\,, \quad \frac{\varphi_v}{M_*} = \frac{\varphi_g}{M_g} - \frac{\varphi_f}{M_f}$
Lagrangian	$+\sqrt{-g}\,\mathcal{L}_g(g,\phi_g) + \sqrt{-f}\,\mathcal{L}_f(f,\phi_f)$	$\sqrt{-\bar{g}}\left[\mathcal{L}^{(2)}_{\text{massless}}(u_{\mu\nu},\varphi_u) + \mathcal{L}^{(2)}_{\text{massive}}(v_{\mu\nu},\varphi_v)\right]$

Massive spin-2 Lagrangian (“Minimal” theory)

- Massive spin-2 action: the same as GW action plus a Fierz-Pauli mass term
- Couples with inflaton perturbations.

$$\mathcal{L}^{(2)}_{vv} = -\frac{1}{2}\nabla_\lambda v_{\mu\nu}\nabla^\lambda v^{\mu\nu} + \nabla_\mu v^{\nu\lambda}\nabla_\nu v^\mu{}_\lambda - \nabla_\mu v^{\mu\nu}\nabla_\nu v + \frac{1}{2}\nabla_\mu v\nabla^\mu v$$
$$+ \left(\bar{R}_{\mu\nu} - M_P^{-2}\,\nabla_\mu\bar{\phi}\nabla_\nu\bar{\phi}\right)\left(v^{\mu\lambda}v_\lambda{}^\nu - \frac{1}{2}v^{\mu\nu}v\right)$$
$$- \frac{1}{2}m^2\left(v^{\mu\nu}v_{\mu\nu} - v^2\right)$$

Identical to GW action

Fierz-Pauli mass term

$$\mathcal{L}^{(2)}_{v\,\varphi_v} = M_P^{-1}\left[\left(\nabla_\mu\bar{\phi}\nabla_\nu\varphi_v + \nabla_\nu\bar{\phi}\nabla_\mu\varphi_v\right)\left(v^{\mu\nu} - \frac{1}{2}\bar{g}^{\mu\nu}v\right) - V'(\bar{\phi})\varphi_v v\right]$$

$$\mathcal{L}^{(2)}_{\varphi_v\varphi_v} = -\frac{1}{2}\nabla_\mu\varphi_v\nabla^\mu\varphi_v - \frac{1}{2}V''(\bar{\phi})\varphi_v^2$$

Spin-2 couples with inflaton perturbations

Specializing to FRW universe

Scalar-Vector-Tensor decomposition

- Decompose the metric perturbation into spatial scalars, vectors and tensor.
- Degrees of freedom: 2 for tensor, 2 for vector, 1 for scalar.
- SVT sectors decouple at quadratic order.

$$v_{00} = a^2 E, \quad v_{0i} = a^2(\partial_i F + G_i), \quad v_{ij} = a^2(\delta_{ij} A + \partial_i \partial_j B + \partial_i C_j + \partial_j C_i + D_{ij})$$

$$\Downarrow$$

$$\mathcal{L}^{(2)}_{\text{massive}} = \mathcal{L}^{(2)}_{\text{tensor}}(D_{ij}, \cdots) + \mathcal{L}^{(2)}_{\text{vector}}(G_i, C_i, \cdots) + \mathcal{L}^{(2)}_{\text{scalar}}(A, B, E, F, \cdots)$$

SVT decomposition (“Minimal” theory)

- Spin-2 dofs: 2 tensor & 2 vector & 1 scalar
- In the scalar sector, the spin-2 dof (B) is mixed with inflaton perturbation

$$\mathcal{L}^{(2)}_{\text{tensor},\mathbf{k}} = \frac{1}{2}a^2\left[\tilde{D}'_{ij}\tilde{D}'_{ij} - (k^2 + a^2m^2)\tilde{D}_{ij}\tilde{D}_{ij}\right]$$

$$\mathcal{L}^{(2)}_{\text{vector},\mathbf{k}} = \frac{a^4k^2m^2}{k^2+a^2m^2}|\tilde{C}'_i|^2 - a^4k^2m^2|\tilde{C}_i|^2$$

$$\mathcal{L}^{(2)}_{\text{scalar},\mathbf{k}} = K_\varphi\,|\tilde{\hat{\varphi}}'_v|^2 - M_\varphi\,|\tilde{\hat{\varphi}}_v|^2 + K_B\,|\tilde{B}'|^2 - M_B\,|\tilde{B}|^2 + L_2\,\tilde{\hat{\varphi}}^{*\prime}_v\tilde{B}' + L_1\,\tilde{\hat{\varphi}}^*_v\tilde{B}' - L_0\,\tilde{\hat{\varphi}}^*_v\tilde{B}$$

Scalar mode coupled with inflaton

$$\tilde{\hat{\varphi}}_v = \tilde{\Pi} + \kappa(\eta)\,\tilde{\mathcal{B}} \qquad \text{and} \qquad \tilde{B} = k^{-2}\tilde{\mathcal{B}}$$

$$\kappa(\eta) = -\frac{L_2}{2k^2K_\varphi}$$

A change of variable decouples the two at early/late times

$$K_\varphi = \frac{a^2}{2}\frac{H^2k^4 + 3a^2\left(m^2 - m_H^2\right)H^2k^2 + \frac{9}{4}a^4m^2\left(m^2 - m_H^2\right)H^2}{H^2k^4 + 3a^2\left(m^2 - m_H^2\right)H^2k^2 + \frac{3}{8}a^4m^2\left(6m^2H^2 - 4H^2m_H^2 - m_H^4\right)} \tag{3.17a}$$

$$M_\varphi = \frac{a^2}{2}\frac{c_{10}k^{10} + c_8k^8 + c_6k^6 + c_4k^4 + c_2k^2 + c_0}{\left[H^2k^4 + 3a^2\left(m^2 - m_H^2\right)H^2k^2 + \frac{3}{8}a^4m^2\left(6m^2H^2 - 4H^2m_H^2 - m_H^4\right)\right]^2} \tag{3.17b}$$

$$\begin{aligned}
c_{10} &= H^4 \\
c_8 &= \tfrac{1}{2}a^2H^2\Big[\left(12m^2H^2 + 8H^4 - 14H^2m_H^2 - m_H^4\right) + 4\frac{HV'(\bar\phi)\bar\phi'}{aM_P^2} + 2H^2V''(\bar\phi)\Big] \\
c_6 &= \tfrac{3}{8}a^4H^2\Big[\left(36m^4H^2 + 72m^2H^4 - 82m^2H^2m_H^2 - 64H^4m_H^2\right. \\
&\qquad \left. - 7m^2m_H^4 + 40H^2m_H^4 + 8m_H^6\right) \\
&\qquad + 8\left(3m^2 - 4m_H^2\right)\frac{HV'(\bar\phi)\bar\phi'}{aM_P^2} \\
&\qquad + 16\left(m^2 - m_H^2\right)H^2V''(\bar\phi)\Big] \\
c_4 &= \tfrac{3}{8}a^6\Big[4H^2\left(9m^6H^2 + 36m^4H^4 + 16m^2H^6 - 30m^4H^2m_H^2 - 76m^2H^4m_H^2\right. \\
&\qquad \left. - 3m^4m_H^4 + 31m^2H^2m_H^4 + 24H^4m_H^4 + 6m^2m_H^6 - 6H^2m_H^6 - 3m_H^8\right) \\
&\qquad - 4m^2H^2\left(H^2 - m_H^2\right)\frac{V'(\bar\phi)^2}{M_P^2} \\
&\qquad + \left(36m^4H^2 + 8m^2H^4 - 94m^2H^2m_H^2 + m^2m_H^4 + 48H^2m_H^4\right)\frac{HV'(\bar\phi)\bar\phi'}{aM_P^2} \\
&\qquad + \left(36m^4H^2 - 58m^2H^2m_H^2 - m^2m_H^4 + 24H^2m_H^4\right)H^2V''(\bar\phi)\Big] \\
c_2 &= \tfrac{9}{32}a^8m^2\Big[H^2\left(18m^6H^2 + 120m^4H^4 + 128m^2H^6 - 78m^4H^2m_H^2 - 384m^2H^4m_H^2\right. \\
&\qquad \left. - 9m^4m_H^4 + 132m^2H^2m_H^4 + 128H^4m_H^4 + 23m^2m_H^6 - 32H^2m_H^6 - 16m_H^8\right) \\
&\qquad - 8H^2\left(2m^2H^2 - 2m^2m_H^2 + m_H^4\right)\frac{V'(\bar\phi)^2}{M_P^2} \\
&\qquad + 4\left(6m^4H^2 - 22m^2H^2m_H^2 + m^2m_H^4 + 14H^2m_H^4\right)\frac{HV'(\bar\phi)\bar\phi'}{aM_P^2} \\
&\qquad + 4\left(m^2 - m_H^2\right)\left(12m^2H^2 - 10H^2m_H^2 - m_H^4\right)H^2V''(\bar\phi)\Big] \\
c_0 &= \tfrac{27}{32}a^{10}m^4\Big[-2H^2\left(2m^2H^2 - 2m^2m_H^2 + m_H^4\right)\frac{V'(\bar\phi)^2}{M_P^2} \\
&\qquad - m^2\left(2H^2 - m_H^2\right)\left(4H^2 + m_H^2\right)\frac{HV'(\bar\phi)\bar\phi'}{aM_P^2} \\
&\qquad + \left(m^2 - m_H^2\right)\left(6m^2H^2 - 4H^2m_H^2 - m_H^4\right)H^2V''(\bar\phi)\Big]
\end{aligned}$$

$$K_B = \frac{a^6m^2}{8}\frac{\left(8m^2H^2 - 6H^2m_H^2 - m^2m_H^2\right)k^4}{H^2k^4 + 3a^2\left(m^2 - m_H^2\right)H^2k^2 + \frac{3}{8}a^4m^2\left(6m^2H^2 - 4H^2m_H^2 - m_H^4\right)} \tag{3.17c}$$

$$M_B = \frac{a^6m^2}{8}\frac{c_{10}k^{10} + c_8k^8 + c_6k^6 + c_4k^4}{\left[H^2k^4 + 3a^2\left(m^2 - m_H^2\right)H^2k^2 + \frac{3}{8}a^4m^2\left(6m^2H^2 - 4H^2m_H^2 - m_H^4\right)\right]^2} \tag{3.17d}$$

$$\begin{aligned}
c_{10} &= H^2\left(8m^2H^2 - 8H^4 - 2H^2m_H^2 - m^2m_H^2\right) \\
c_8 &= a^2H^2\Big[\left(30m^4H^2 + 32m^2H^4 - 96H^6 - 3m^4m_H^2 - 56m^2H^2m_H^2\right. \\
&\qquad \left. + 48H^4m_H^2 + 5m^2m_H^4 + 6H^2m_H^4\right) \\
&\qquad + \left(4m^2 - 24H^2\right)\frac{HV'(\bar\phi)\bar\phi'}{aM_P^2}\Big] \\
c_6 &= \tfrac{3}{8}a^4m^2\Big[\left(96m^4H^4 + 144m^2H^6 - 6m^4H^2m_H^2 - 252m^2H^4m_H^2 - 192H^6m_H^2\right. \\
&\qquad \left. + 8m^2H^2m_H^4 + 200H^4m_H^4 - 10H^2m_H^6 - m^2m_H^6\right) \\
&\qquad + \left(8m^2m_H^2 - 16H^2m_H^2\right)\frac{HV'(\bar\phi)\bar\phi'}{aM_P^2}\Big] \\
c_4 &= \tfrac{3}{8}a^6m^4\Big[\left(36m^4H^4 - 48m^2H^6 + 64H^8 - 12m^2H^4m_H^2 - 32H^6m_H^2\right. \\
&\qquad \left. - 12m^2H^2m_H^4 + 4H^4m_H^4 + 12H^2m_H^6 - 3m^2m_H^6 + 2m_H^8\right) \\
&\qquad - \left(24m^2H^2 - 16H^4 - 12m^2m_H^2 - 8H^2m_H^2 + 8m_H^4\right)\frac{HV'(\bar\phi)\bar\phi'}{aM_P^2}\Big]
\end{aligned}$$

$$L_2 = \frac{a^3m^2\bar\phi'}{2M_PH}\frac{H^2k^4 + \frac{3}{2}a^2\left(m^2 - m_H^2\right)H^2k^2}{H^2k^4 + 3a^2\left(m^2 - m_H^2\right)H^2k^2 + \frac{3}{8}a^4m^2\left(6m^2H^2 - 4H^2m_H^2 - m_H^4\right)} \tag{3.17e}$$

$$L_1 = -\frac{a^4m^2\bar\phi'}{M_P}\frac{\left(H^2 - \frac{1}{4}m_H^2 - \frac{1}{2}\frac{aHV'(\bar\phi)}{\phi'}\right)k^4 - \frac{3}{2}a^2\left(m^2 - m_H^2\right)\left(H^2 + \frac{1}{4}m_H^2 + \frac{1}{2}\frac{aHV'(\bar\phi)}{\phi'}\right)k^2}{H^2k^4 + 3a^2\left(m^2 - m_H^2\right)H^2k^2 + \frac{3}{8}a^4m^2\left(6m^2H^2 - 4H^2m_H^2 - m_H^4\right)} \tag{3.17f}$$

$$L_0 = \frac{a^3m^2\bar\phi'}{2M_PH}\frac{c_{10}k^{10} + c_8k^8 + c_6k^6 + c_4k^4 + c_2k^2}{\left[H^2k^4 + 3a^2\left(m^2 - m_H^2\right)H^2k^2 + \frac{3}{8}a^4m^2\left(6m^2H^2 - 4H^2m_H^2 - m_H^4\right)\right]^2} \tag{3.17g}$$

$$\begin{aligned}
c_{10} &= H^4 \\
c_8 &= \tfrac{1}{2}a^2H^4\Big[\left(9m^2 + 12H^2 - 13m_H^2\right) - 4\frac{aHV'(\bar\phi)}{\phi'}\Big] \\
c_6 &= \tfrac{3}{8}a^4H^2\Big[\left(18m^4H^2 + 32m^2H^4 + 64H^6 - 48m^2H^2m_H^2 - 64H^4m_H^2\right. \\
&\qquad \left. + m^2m_H^4 + 28H^2m_H^4\right) \\
&\qquad + 8\left(-4m^2H^2 + 4H^4 + m^2m_H^2\right)\frac{aHV'(\bar\phi)}{\bar\phi'}\Big] \\
c_4 &= \tfrac{3}{16}a^6m^2H^2\Big[\left(18m^4H^2 - 24m^2H^4 + 256H^6 - 54m^2H^2m_H^2 - 160H^4m_H^2\right. \\
&\qquad \left. + 9m^2m_H^4 + 60H^2m_H^4 - 7m_H^6\right) \\
&\qquad + 4\left(-30m^2H^2 + 32H^4 + 12m^2m_H^2 + 4H^2m_H^2 - 7m_H^4\right)\frac{aHV'(\bar\phi)}{\bar\phi'}\Big] \\
c_2 &= \tfrac{9}{16}a^8m^4H^2\left(2H^2 - m_H^2\right)\Big[-\left(4H^2 + m_H^2\right)\left(3m^2 - 4H^2 - m_H^2\right) \\
&\qquad + 4\left(-3m^2 + 2H^2 + 2m_H^2\right)\frac{aHV'(\bar\phi)}{\phi'}\Big]
\end{aligned}$$

Generalized Higuchi bound (“Minimal” theory)

- The theory has a lower bound on mass m, below which the scalar sector constains a ghost with wrong sign kinetic term
- Higuchi: derived the bound for de-Sitter (the Higuchi bound)
- This work: derived the bound for FRW

$$m^2 > m_H^2(\eta) = 2H^2 + 2a^{-1}H' = 2H(\eta)^2\left[1-\epsilon\right]$$

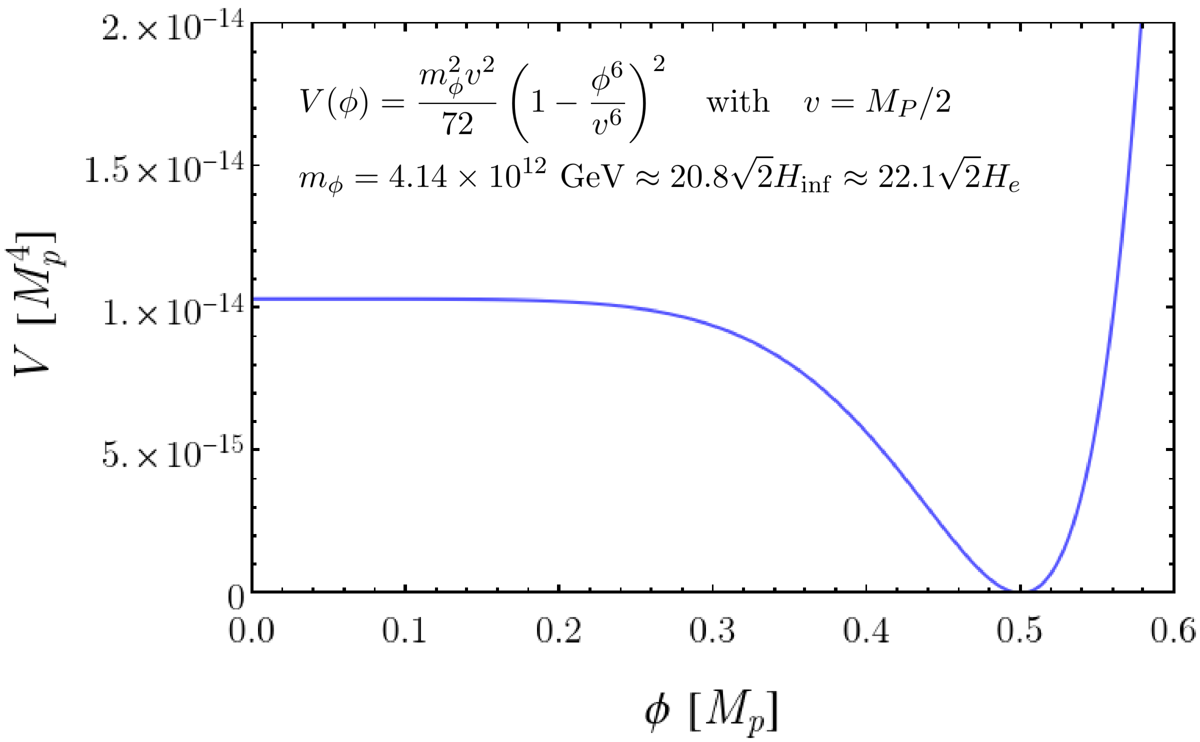
$$L_{S,\boldsymbol{k}} = K_\varphi\,|\tilde{\hat{\varphi}}_v'|^2 - M_\varphi\,|\tilde{\hat{\varphi}}_v|^2 + K_B\,|\tilde{B}'|^2 - M_B\,|\tilde{B}|^2 + L_2\,\tilde{\hat{\varphi}}_v^{*\prime}\tilde{B}' + L_1\,\tilde{\hat{\varphi}}_v^{*}\tilde{B}' - L_0\,\tilde{\hat{\varphi}}_v^{*}\tilde{B}$$

A. Higuchi (1987), *Forbidden mass range for spin-2 field theory in de Sitter spacetime*

Results / Constraints

Hilltop inflation

- Use hilltop inflation model to solve for cosmic expansion history a(eta).
- The massive spin-2 field is a spectator on this background.



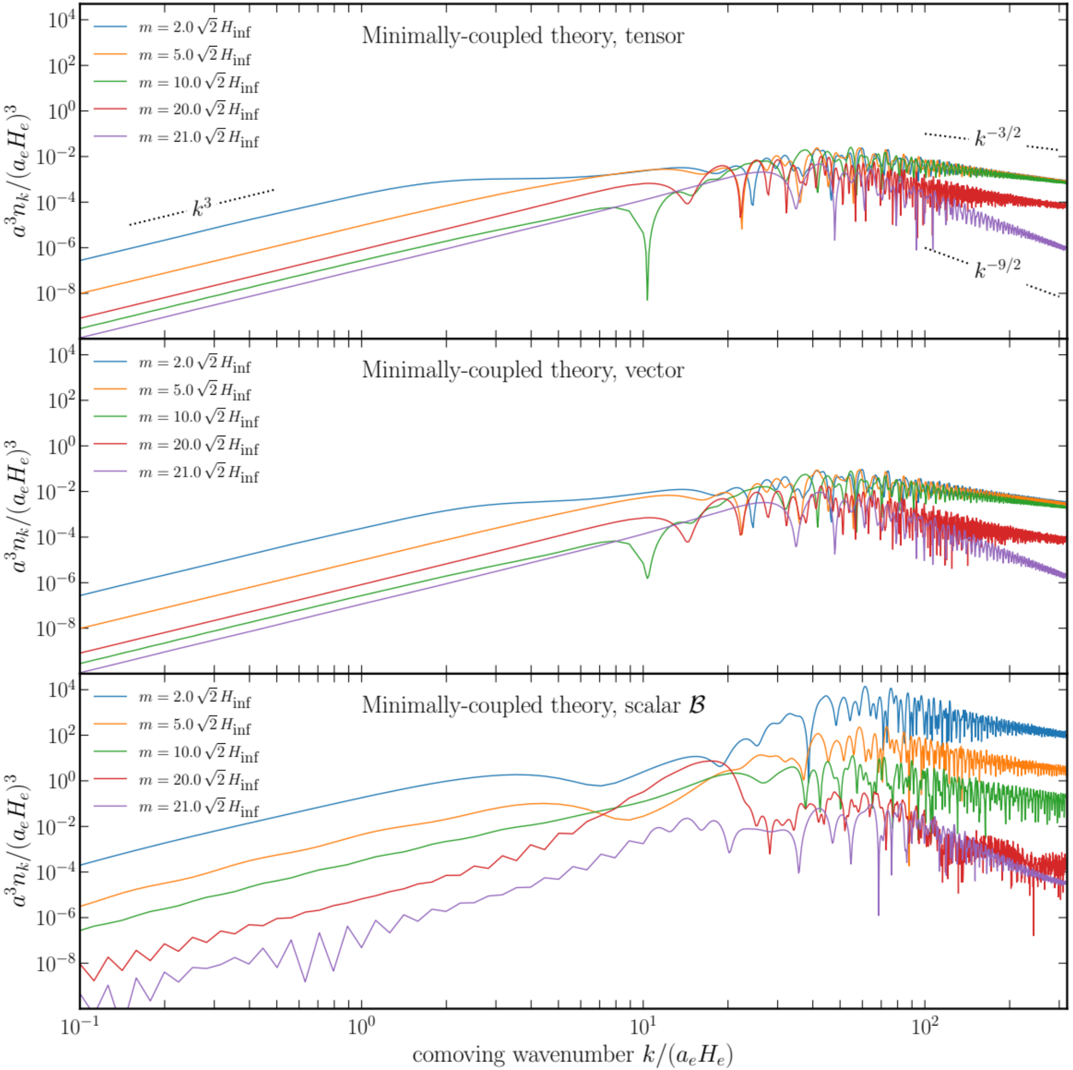
Number density
("Minimal" theory)

- The number density spectrum:
$$n_k(\eta) = a(\eta)^{-3}\frac{k^3}{2\pi^2}|\beta_k|^2$$
- Low-k modes driven by superhorizon dynamics.

$$\omega_k^2 \approx (1-\delta)k^2 + a(\eta)^2 H_{\rm inf}^2(\mu^2+\alpha)$$

$$n_k \propto k^{3-2\nu} \quad \text{where} \quad \nu = \sqrt{\frac{1}{4} - \alpha - \frac{m^2}{H_{\rm inf}^2}}$$

$$n_k \propto k^3 \quad \text{if} \quad \nu \quad \text{imaginary}$$



Chung, Kolb, Riotto & Senatore (2004).

Number density
("Minimal" theory)

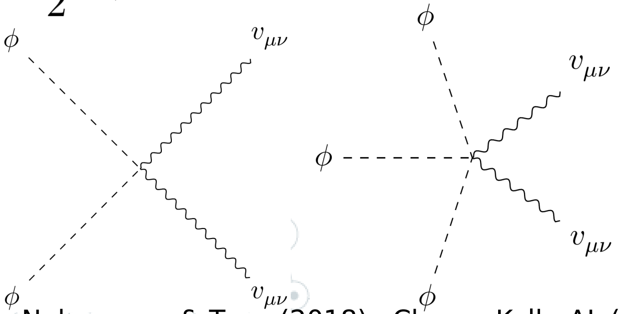
- The number density spectrum:

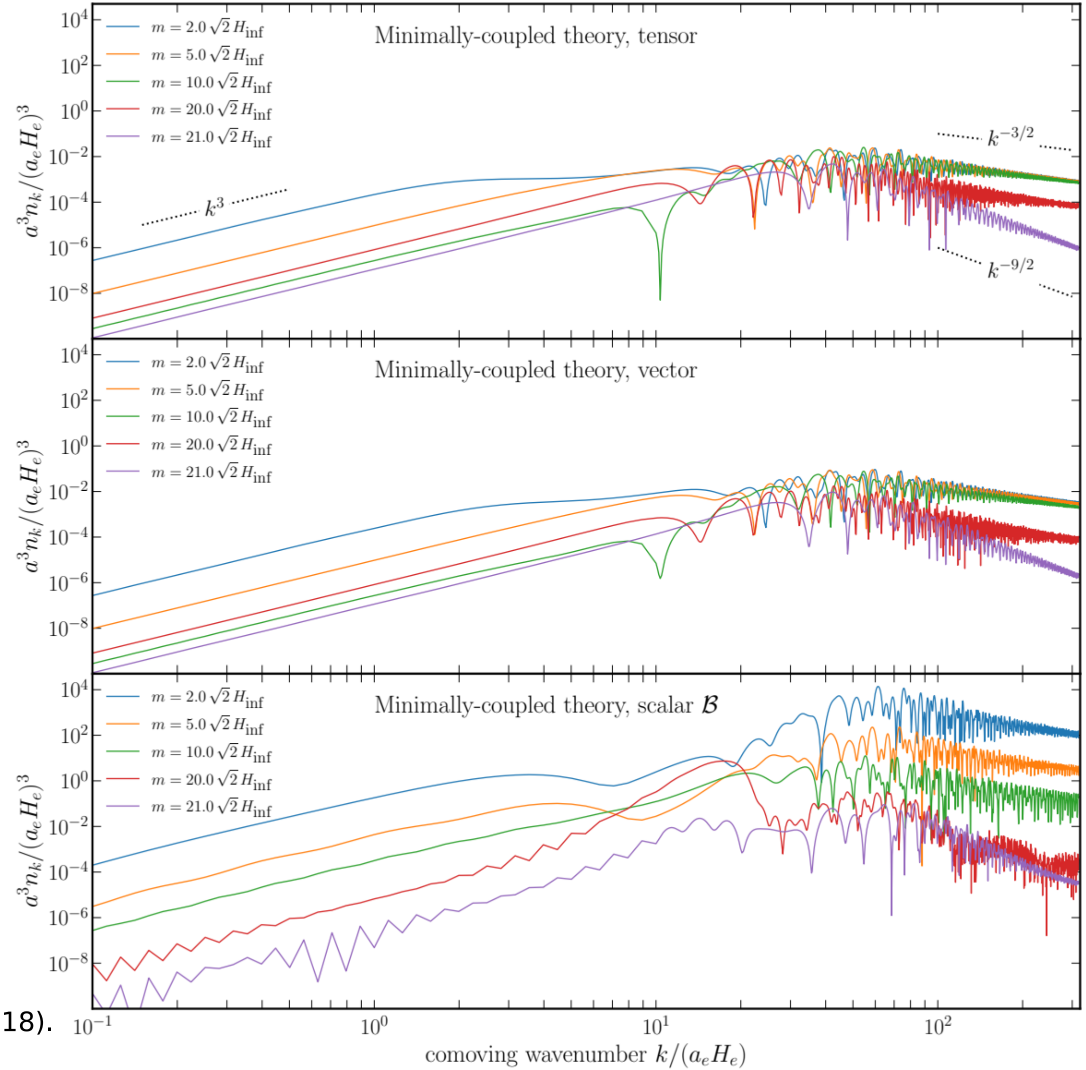
$$n_k(\eta) = a(\eta)^{-3} \frac{k^3}{2\pi^2} |\beta_k|^2$$

- High-k modes driven by homogeneous oscillation of inflaton (after inflation).

$$m < m_\phi \Rightarrow \phi\phi \to \chi\chi \Rightarrow n_k \propto k^{-3/2}$$

$$m < \frac{3}{2} m_\phi \Rightarrow \phi\phi\phi \to \chi\chi \Rightarrow n_k \propto k^{-9/2}$$



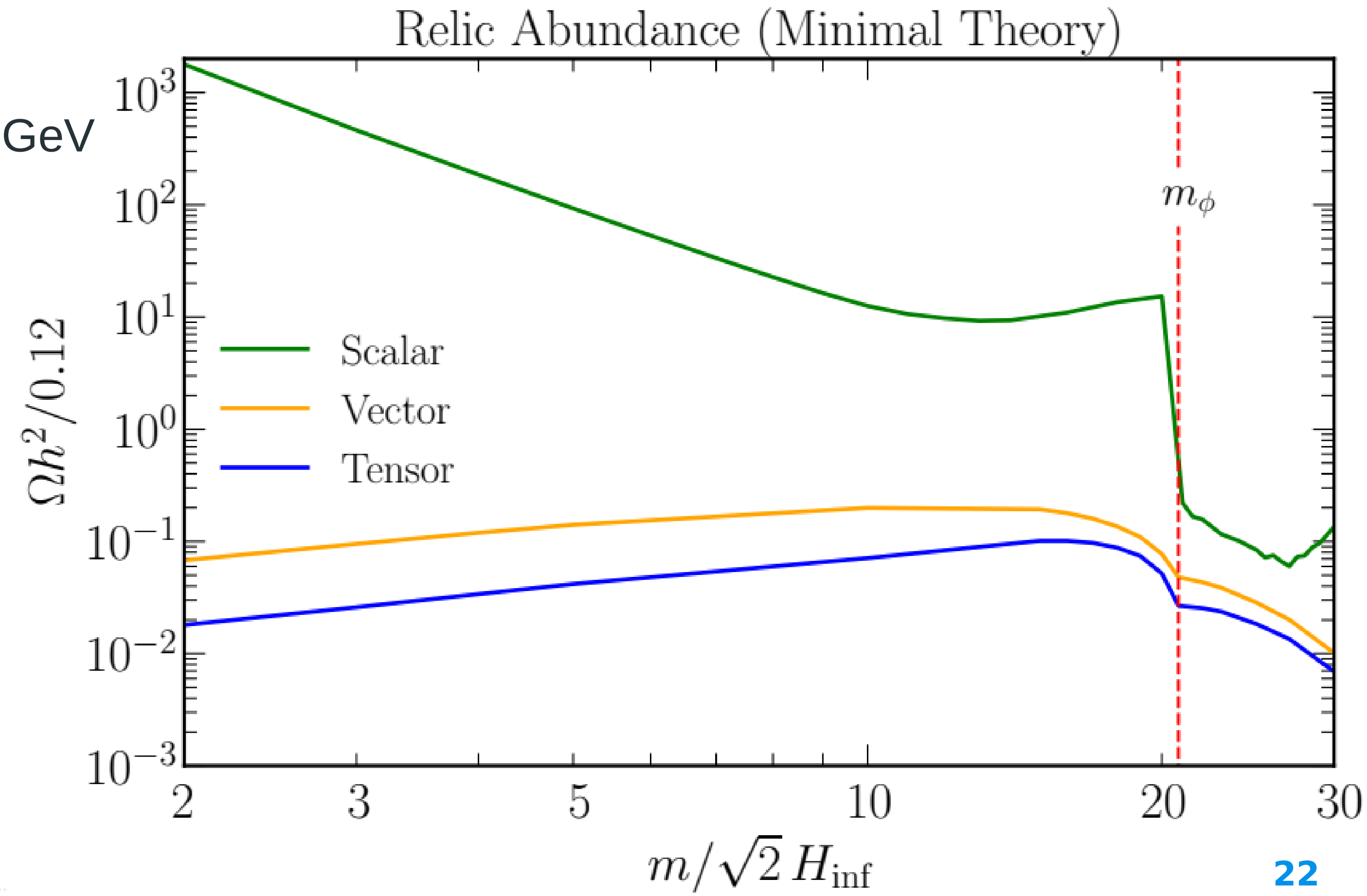


Ema, Nakayama, & Tang (2018). Chung, Kolb, AL (2018). Basso, Chung, Kolb, Long (2022).

Relic abundance

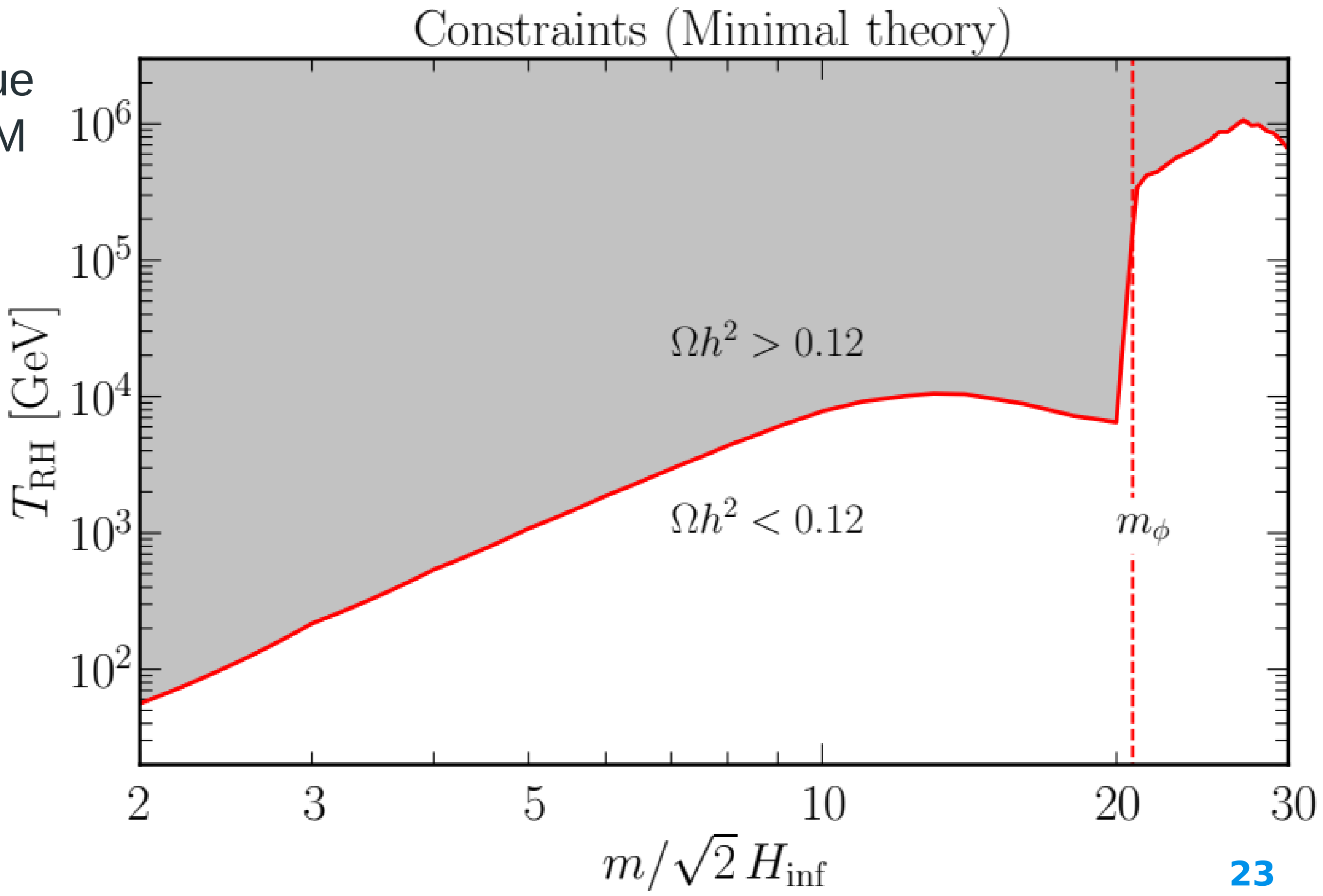
(“Minimal” theory)

- Assuming reheating temperature of 10^5 GeV
- Sharp drop at red-dashed line where ϕϕ -> χχ scattering channel becomes kinetically forbidden



Constraints
(“Minimal” theory)

- Gray zone excluded due to overproduction of DM



Summary

- Superheavy spin-2 particles can be gravitationally produced during inflation.
- CGPP is a valid production mechanism for massive spin-2 DM.
- Aside: We find a generalized Higuchi bound for the massive spin-2 in FRW.

Extra slides

Inflation: where GPP happens

- Why inflation:
  - Significant particle production due to rapid change in scale factor.
  - We know the field's initial condition during inflation (aka Bunch-Davies initial condition).

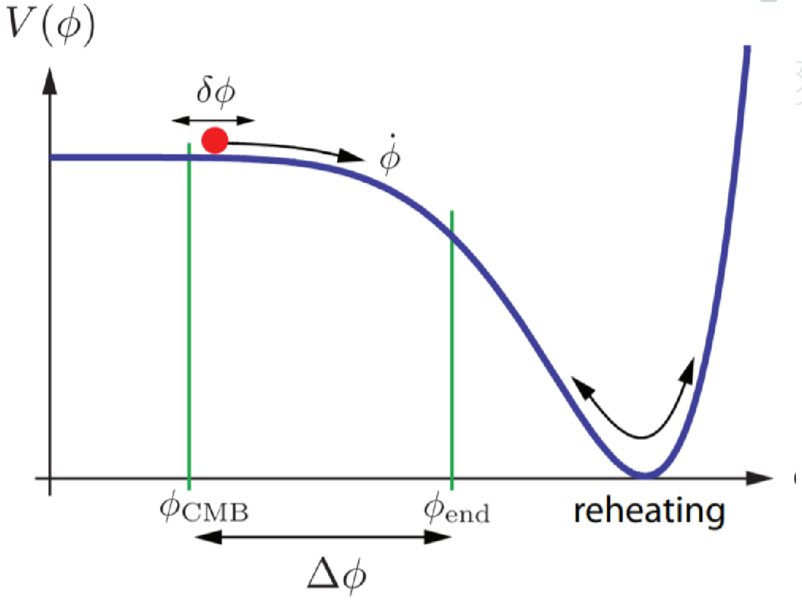


Figure due to Anupam Mazumdar

The "Nonminimal" theory

- The inflaton couples the "effective" metric.
- The "massive" sector only contain the massive spin-2 field.

	Original field variables	Mass eigenstates
Metric	$g_{\mu\nu} = \bar{g}_{\mu\nu} + \frac{2}{M_g} h_{\mu\nu}\,, \quad f_{\mu\nu} = \bar{g}_{\mu\nu} + \frac{2}{M_f} k_{\mu\nu}$	$\frac{u_{\mu\nu}}{M_*} = \frac{h_{\mu\nu}}{M_f} + \frac{k_{\mu\nu}}{M_g}\,, \quad \frac{v_{\mu\nu}}{M_*} = \frac{h_{\mu\nu}}{M_g} - \frac{k_{\mu\nu}}{M_f}$
Inflaton	$\phi_\star = \bar{\phi}_\star + \varphi_\star$	
Lagrangian	$+\sqrt{-g_\star}\,\mathcal{L}_\star(g_\star, \phi_\star)$	$\sqrt{-\bar{g}}\left[\mathcal{L}^{(2)}_{\text{massless}}(u_{\mu\nu}, \varphi_\star) + \mathcal{L}^{(2)}_{\text{massive}}(v_{\mu\nu})\right]$

$$(g_\star)_{\mu\nu} = \frac{a^2}{(a+b)^2} g_{\mu\nu} + \frac{ab}{(a+b)^2}\left(g_{\mu\lambda}\left(\sqrt{g^{-1}f}\right)^\lambda{}_\nu + \left(\sqrt{g^{-1}f}\right)_\mu{}^\lambda g_{\lambda\nu}\right) + \frac{b^2}{(a+b)^2} f_{\mu\nu}$$

Massive spin-2 Lagrangian (“Nonminimal” theory)

- Massive spin-2 action: contains a Fierz-Pauli mass term
- Contains terms that depend on ϕ potential. Gives “wrong” de-Sitter limit.

$$
\begin{aligned}
\mathcal{L}^{(2)}_{vv} = &-\tfrac{1}{2}\nabla_\lambda v_{\mu\nu}\nabla^\lambda v^{\mu\nu} + \nabla_\mu v^{\nu\lambda}\nabla_\nu v^\mu{}_\lambda - \nabla_\mu v^{\mu\nu}\nabla_\nu v + \tfrac{1}{2}\nabla_\mu v\nabla^\mu v \\
&+ \left(\bar{R}_{\mu\nu} + \tfrac{1}{2}M_P^{-2}\left(\nabla_\mu\bar{\phi}\nabla_\nu\bar{\phi} + \bar{g}_{\mu\nu}\bar{\mathcal{L}}(\bar{g},\bar{\phi})\right)\right)v^{\mu\lambda}v_\lambda{}^\nu \\
&- \tfrac{1}{2}\left(\bar{R}_{\mu\nu} + M_P^{-2}\left(\nabla_\mu\bar{\phi}\nabla_\nu\bar{\phi} + \bar{g}_{\mu\nu}\bar{\mathcal{L}}(\bar{g},\bar{\phi})\right)\right)v^{\mu\nu}v \\
&- \tfrac{1}{2}m^2\left(v_{\mu\nu}v^{\mu\nu} - v^2\right)
\end{aligned}
$$

Distinct from GW action

Fierz-Pauli mass term

SVT decomposition ("Nonminimal" theory)

- Spin-2 dofs: 2 tensor & 2 vector & 1 scalar
- The scalar sector now contains only one d.o.f.

$$\mathcal{L}^{(2)}_{\mathrm{tensor},\mathbf{k}} = \frac{1}{2}a^2\left[\tilde{D}'_{ij}\tilde{D}'_{ij} - \left(k^2 + a^2\mu_2^2\right)\tilde{D}_{ij}\tilde{D}_{ij}\right]$$

$$\mathcal{L}^{(2)}_{\mathrm{vector},\mathbf{k}} = \frac{a^4k^2\mu_1^2}{k^2 + a^2\mu_1^2}|\tilde{C}'_i|^2 - a^4k^2\mu_2^2|\tilde{C}_i|^2$$

$$\mathcal{L}^{(2)}_{\mathrm{scalar},\mathbf{k}} = K_B\,|\tilde{B}'|^2 - M_B\,|\tilde{B}|^2$$

$$\mu_1^2 = m^2 - \Lambda + 3H^2 - a^{-1}H'$$

$$\mu_2^2 = m^2 - \Lambda + 3H^2 + 2a^{-1}H'$$

SVT

- Spin-2 dof
- In the scal

$$\mathcal{L}^{(2)}_{\text{tensor},\mathbf{k}} = \frac{1}{2}$$

$$\mathcal{L}^{(2)}_{\text{vector},\mathbf{k}} = \frac{}{k}$$

$$\mathcal{L}^{(2)}_{\text{scalar},\mathbf{k}} = K$$

$$\mu_1^2 = m^2$$

$$\mu_2^2 = m^2$$

$$K_B = \frac{a^4}{P}\left[c_6 k^6 + c_4 k^4\right] \tag{3.46a}$$

$$c_6 = -4\dot{H}^2$$

$$c_4 = 3a^2\left(m^2+H^2\right)\left(m^2+3H^2-\dot{H}\right)\left(m^2+3H^2+2\dot{H}\right)$$

$$M_B = \frac{a^6}{P^2}\left[c_{10}k^{10} + c_8 k^8 + c_6 k^6 + c_4 k^4\right] \tag{3.46b}$$

$$\begin{aligned} c_{10} = {} & 12\left(m^2+H^2\right)\left(m^2+3H^2\right)^3 + 16\left(m^2+3H^2\right)^2\left(6m^2+7H^2\right)\dot{H} \\ & + 4\left(m^2+3H^2\right)\left(63m^2+71H^2\right)\dot{H}^2 + 8\left(25m^2+27H^2\right)\dot{H}^3 - 48\dot{H}^4 \\ & - 32H\left(m^2+3H^2\right)\dot{H}\ddot{H} - 48H\dot{H}^2\ddot{H} \end{aligned}$$

$$\begin{aligned} c_8 = {} & 12a^2\left(m^2+3H^2+2\dot{H}\right) \times \Big[2\left(m^2+H^2\right)\left(m^2+3H^2\right)^2\left(2m^2+5H^2\right) \\ & + \left(m^2+3H^2\right)\left(19m^4+64m^2H^2+49H^4\right)\dot{H} + 2\left(7m^4+20m^2H^2+17H^4\right)\dot{H}^2 \\ & - \left(23m^2+25H^2\right)\dot{H}^3 + 2\dot{H}^4 - 2H\left(m^2+H^2\right)\left(m^2+3H^2\right)\ddot{H} \\ & - 4H\left(m^2+H^2\right)\dot{H}\ddot{H}\Big] \end{aligned}$$

$$\begin{aligned} c_6 = {} & 9a^4\left(m^2+H^2\right)\left(m^2+3H^2-\dot{H}\right)\left(m^2+3H^2+2\dot{H}\right)^2 \\ & \times \left[7\left(m^2+H^2\right)\left(m^2+3H^2\right) + 17\left(m^2+H^2\right)\dot{H} - 8\dot{H}^2\right] \end{aligned}$$

$$c_4 = 27a^6\left(m^2+H^2\right)^2\left(m^2+3H^2-\dot{H}\right)^2\left(m^2+3H^2+2\dot{H}\right)^3$$

where $\dot{H} \equiv a^{-1}H'$, $\ddot{H} \equiv a^{-2}H'' - a^{-1}HH'$, and

$$\begin{aligned} P = {} & 4\left[m^2+3H^2+3\dot{H}\right]k^4 \\ & + 12a^2\left[\left(m^2+H^2\right)\left(m^2+3H^2\right) + 2\left(m^2+H^2\right)\dot{H} - \dot{H}^2\right]k^2 \\ & + 9a^4\left(m^2+H^2\right)\left(m^2+3H^2-\dot{H}\right)\left(m^2+3H^2+2\dot{H}\right). \end{aligned} \tag{3.47}$$

Gradient instability ("Nonminimal" theory, vector sector)

- If the squared effective sound speed (c_s^2) is negative at any time, the mode function blows up at high-k due to a period of exponential growth.
- This scenario usually happens around end of inflation, when Hubble-prime is large enough.

$$\tilde{\chi}_r'' \approx -c_s^2 k^2 \tilde{\chi}_r \quad \Rightarrow \quad \tilde{\chi}_r \propto \exp\left[\pm \int^{\eta} \mathrm{d}\eta' \, k\sqrt{-c_s^2}\right]$$

$$c_s^2 \propto m^2 - \Lambda + 3H^2 + 2a^{-1}H'$$

Ghost instability ("Nonminimal" theory, scalar sector)

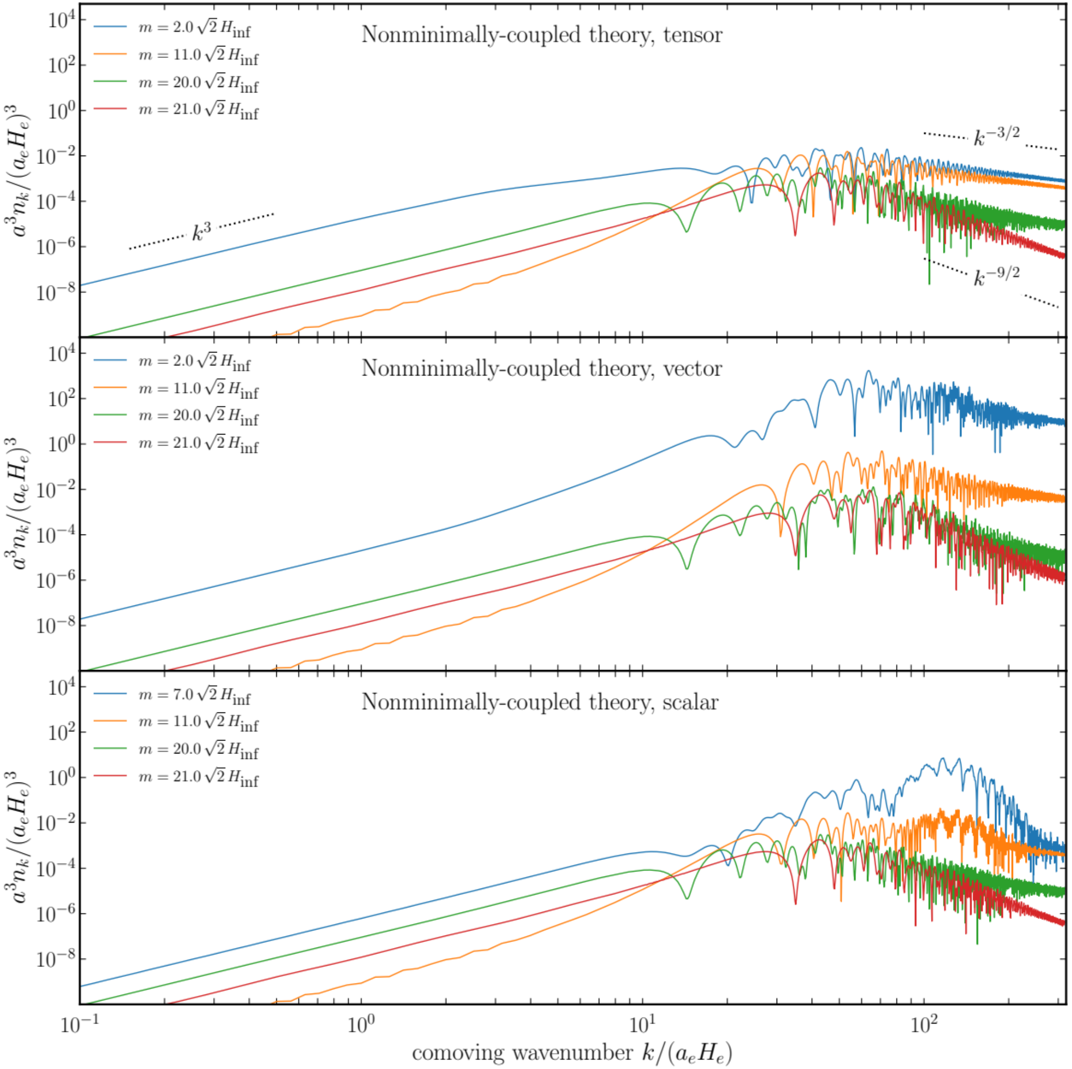
- The kinetic term can momentarily vanish for small enough mass.
- Effectively sets a UV cutoff p_max for physical momentum:

$$p_{\max}(\eta) = \sqrt{\frac{3}{4}} H \frac{\sqrt{m^2/H^2+1}\,\sqrt{m^2/H^2+3+\epsilon}\,\sqrt{m^2/H^2+3-2\epsilon}}{|\epsilon|}$$

Number density
(“Nonminimal” theory)

- Low-k modes have k^3 power law. (driven by superhorizon dynamics)
- High-k modes have k^(-3/2) or k^(-9/2) power law depending on mass of spin-2 particle. (driven by ϕϕ -> χχ or ϕϕϕ -> χχ scattering channels after inflation)

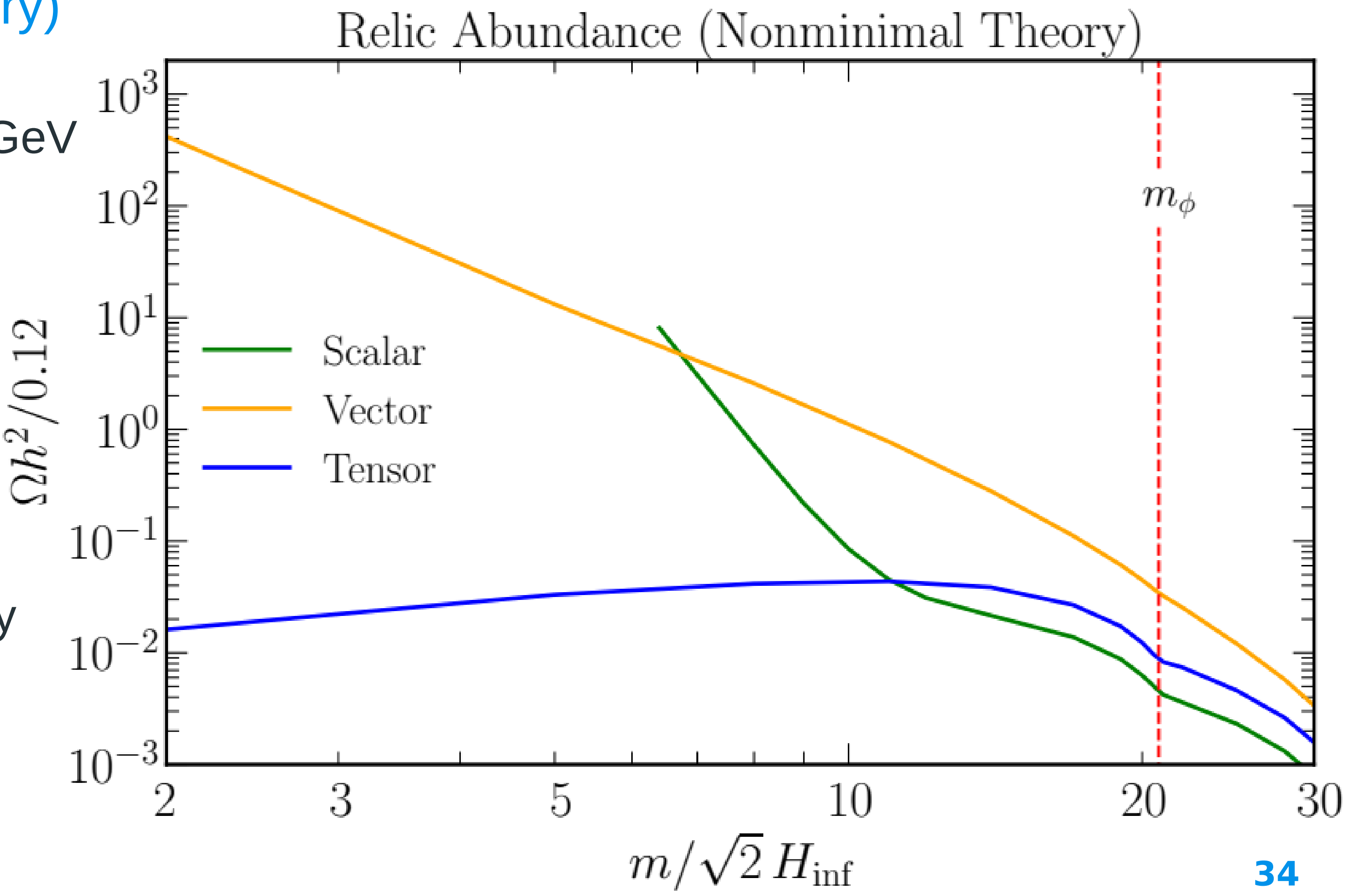
$$n_k(\eta) = a(\eta)^{-3} \frac{k^3}{2\pi^2} |\beta_k|^2$$



Relic abundance

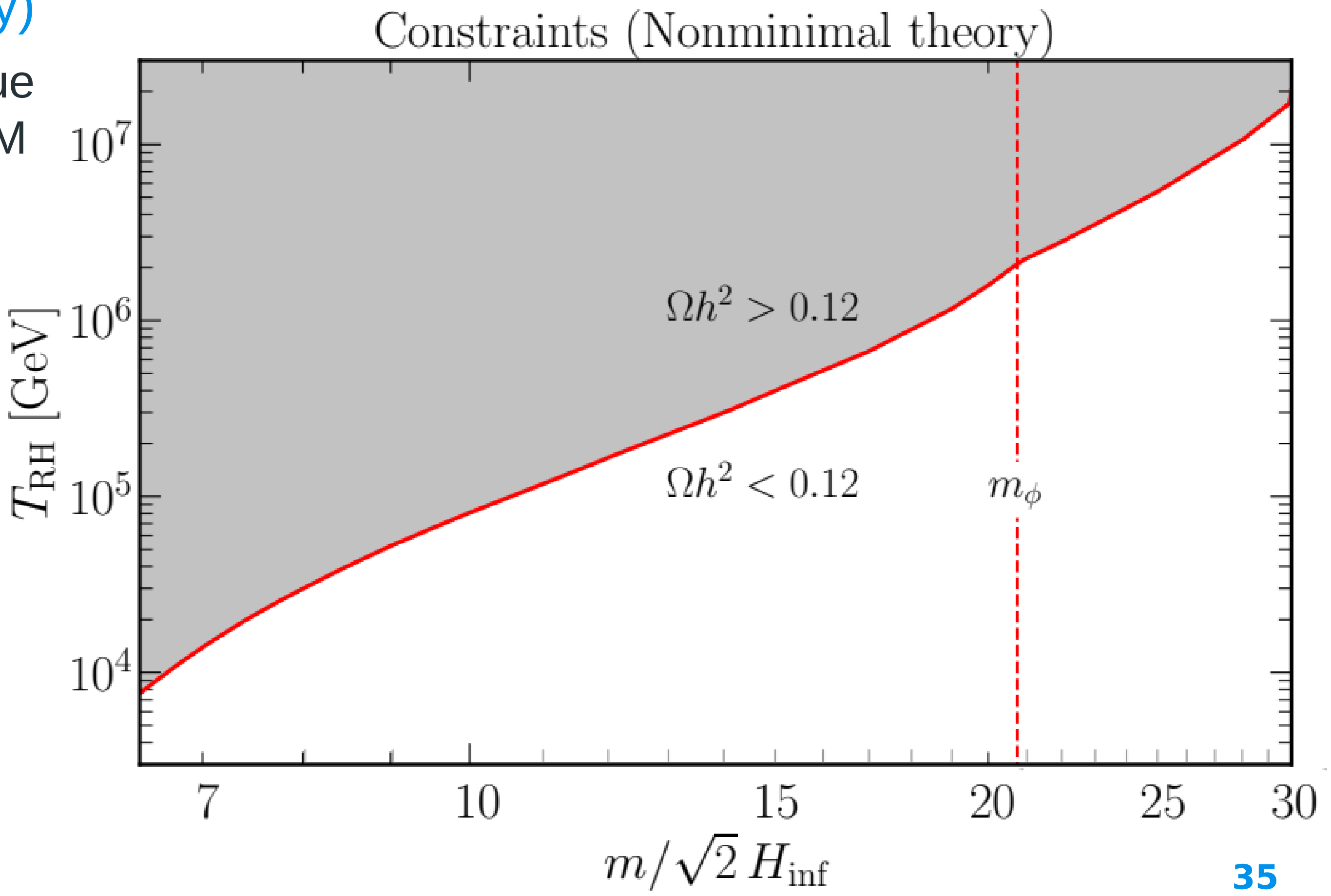
(“Nonminimal” theory)

- Assuming reheating temperature of 10^5 GeV
- Sharp drop at red-dashed line because ϕϕ -> χχ scattering channel becomes kinetically forbidden
- No result for m < 7√2 H_inf in scalar sector due to ghost instability



Constraints
(“Nonminimal” theory)

- Gray zone excluded due to overproduction of DM



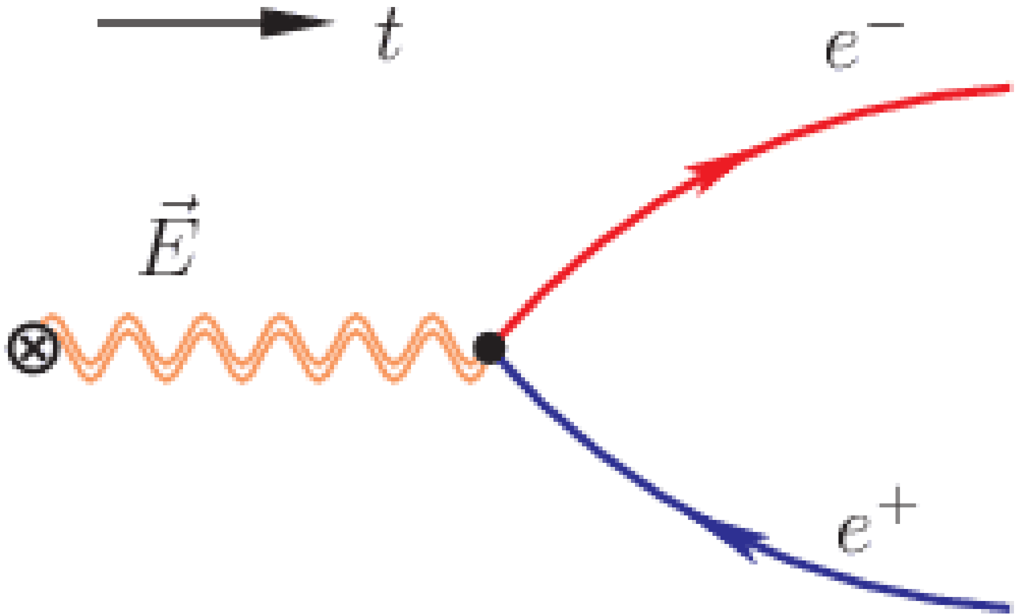
$$L_{S,\boldsymbol{k}} = \frac{1}{2}\left[|\tilde{\chi}'_{\Pi}|^2 - \left(k^2 + a^2 V''(\bar{\phi})\right)|\tilde{\chi}_{\Pi}|^2\right] + \frac{1}{2}\left[|\tilde{\chi}'_{\mathcal{B}}|^2 - \left(k^2 + a^2 m^2\right)|\tilde{\chi}_{\mathcal{B}}|^2\right] + \mathcal{O}(H/m)$$

Formulas

$$\Omega h^2 \approx 0.12 \left(\frac{m}{10^{10}\ \mathrm{GeV}}\right)\left(\frac{H_e}{10^{10}\ \mathrm{GeV}}\right)\left(\frac{T_{\mathrm{RH}}}{10^{8}\ \mathrm{GeV}}\right)\left(\frac{a^3 n}{a_e^3 H_e^3}\right)$$

$$V(\bar{\phi}) = \frac{m_\phi^2 v^2}{72}\left(1 - \frac{\bar{\phi}^6}{v^6}\right)^2$$

$$\lim_{\eta \to -\infty} \tilde{\chi}(\eta, \boldsymbol{k}) = \frac{1}{\sqrt{2k}} e^{-ik\eta}$$


t
$\vec{E}$
e^-
e^+