Probing ultralight dark matter with gravity wave detectors

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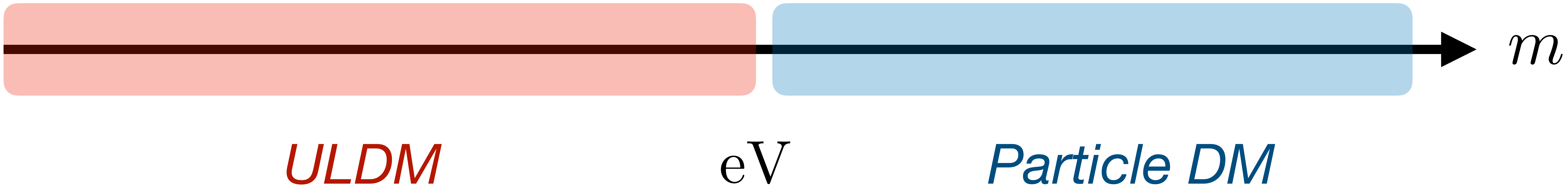
Ultralight Dark Matter

we define *ultralight dark matter (ULDM)*
as *bosonic DM candidates with* $\mathrm{m} < \mathrm{eV}$

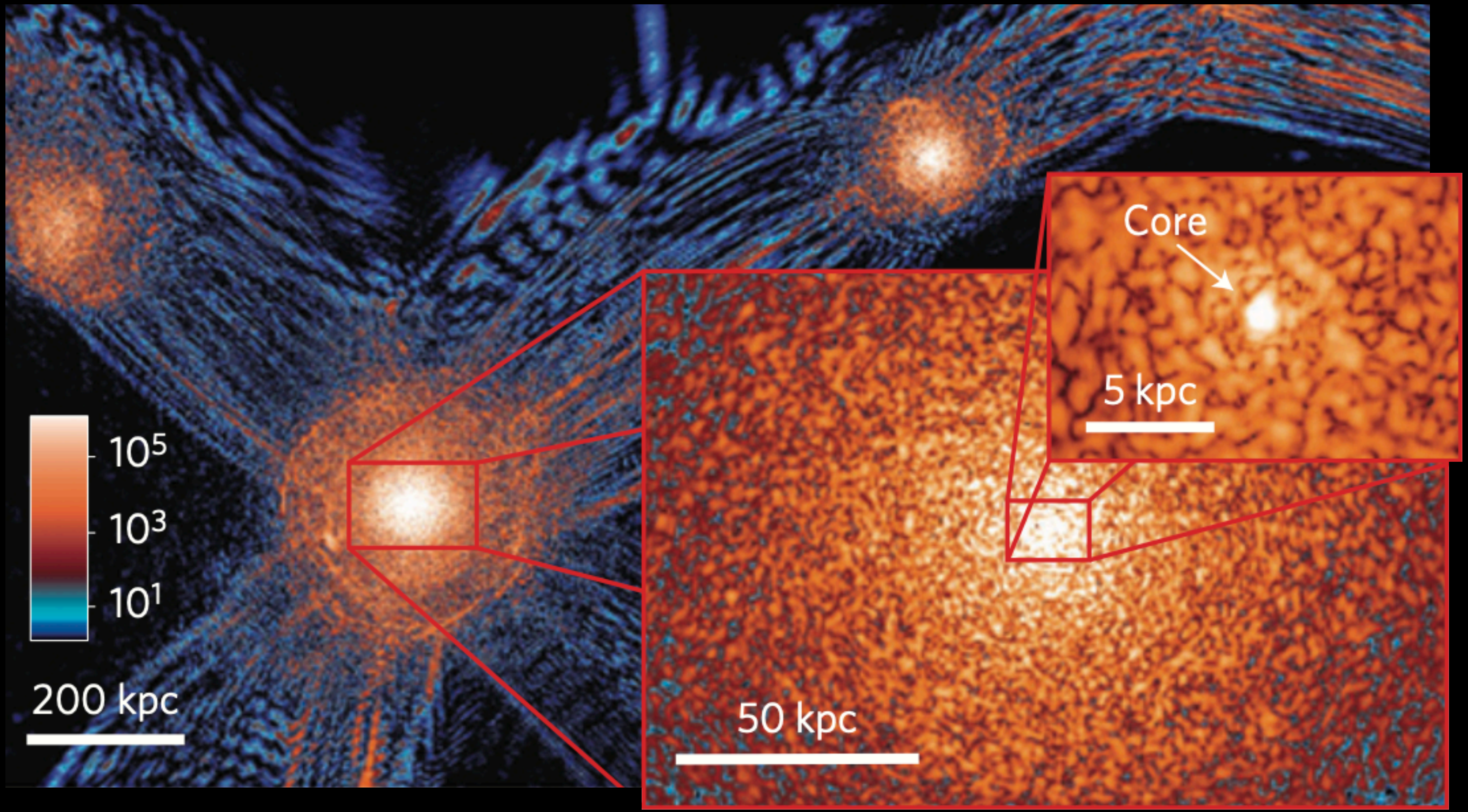


we define *ultralight dark matter (ULDM)*
as *bosonic DM candidates with* $\mathrm{m} < \mathrm{eV}$

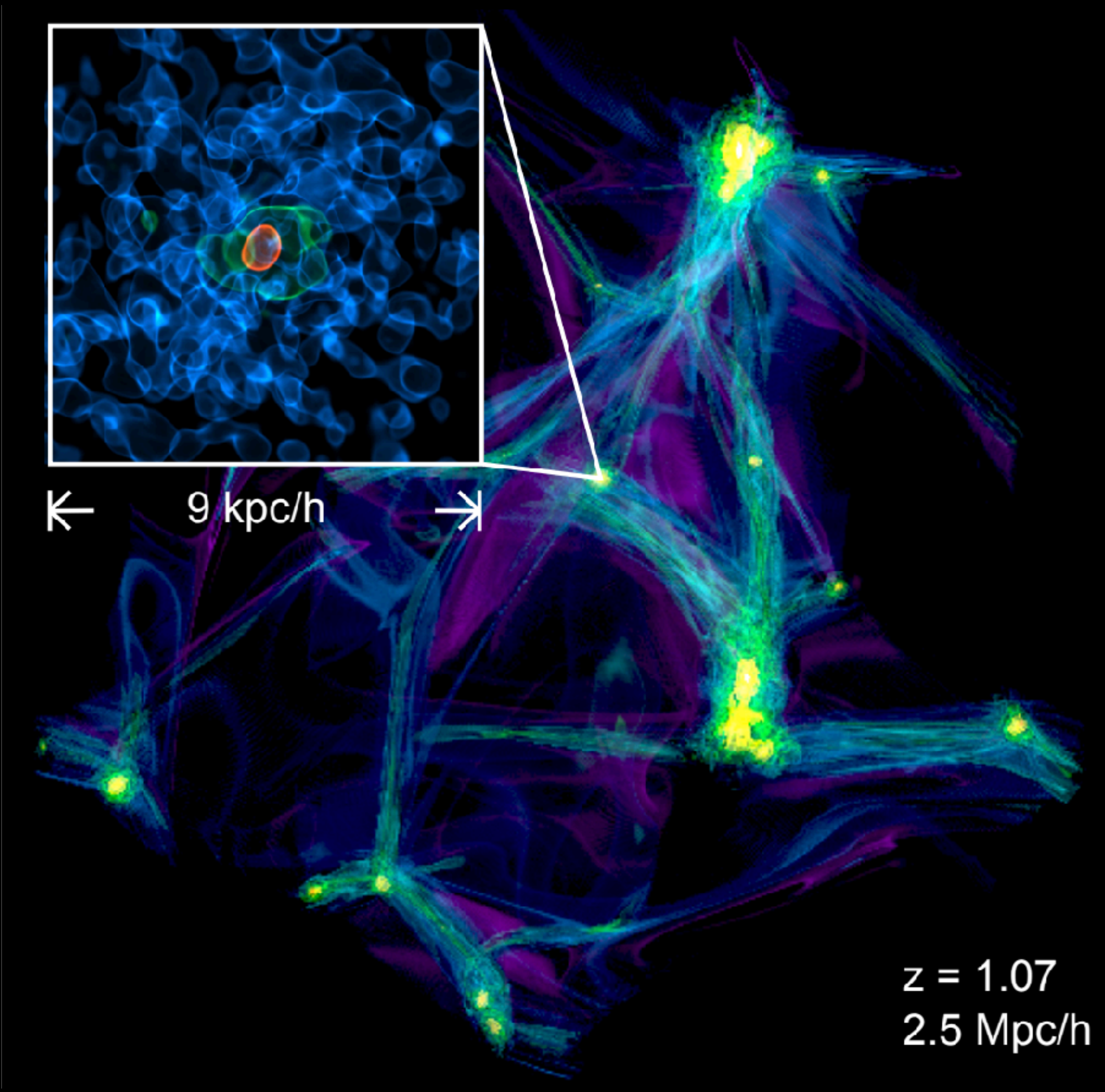
$$m \lesssim 10\,\mathrm{eV}$$

$$N_{\rm occ} \sim n_{\rm dm}\lambda^3 \sim \left(\frac{10\,\mathrm{eV}}{m}\right)^4$$

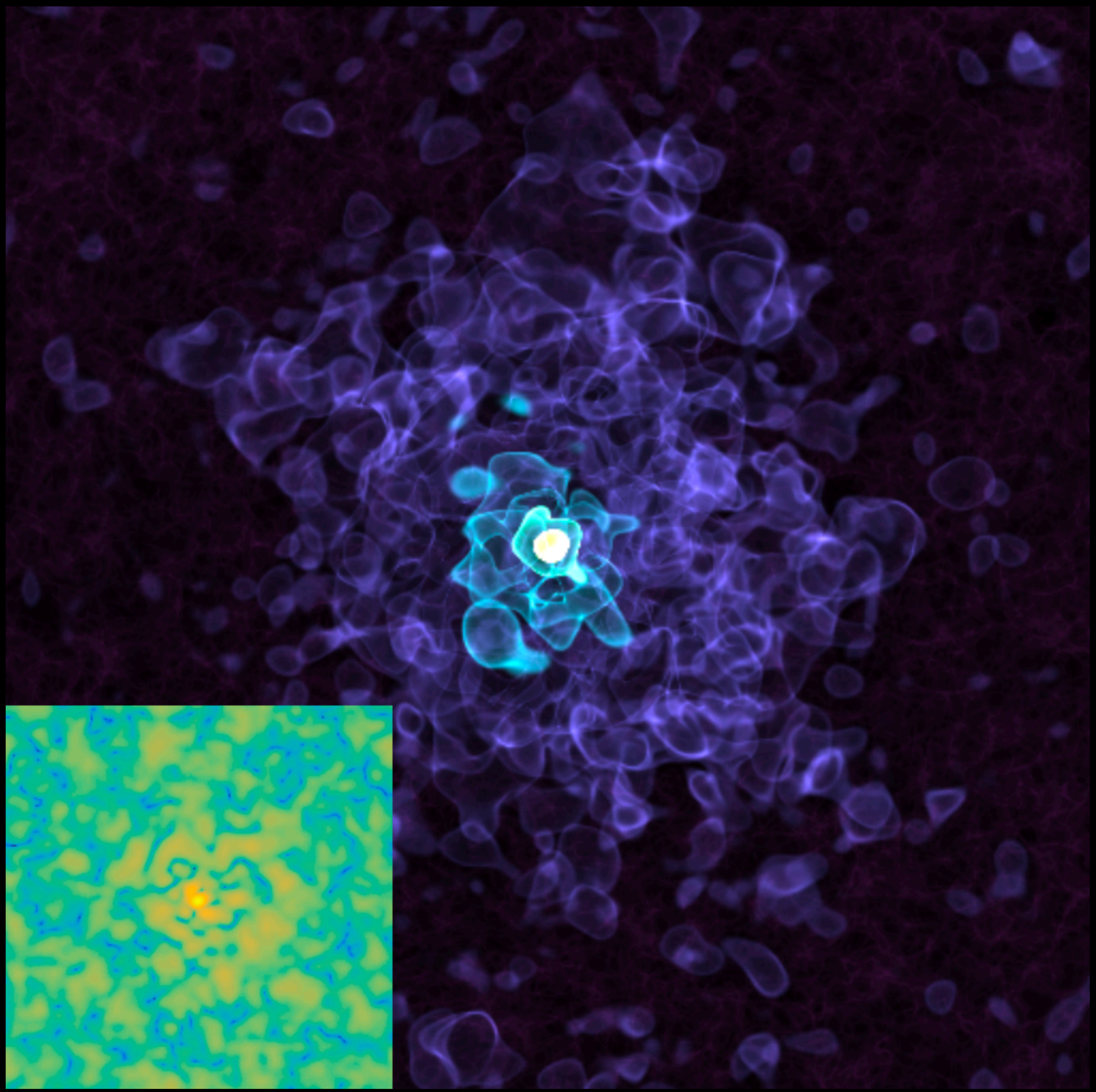
Schive, Chiueh, Broadhurst (14)



Mocz et al (17)

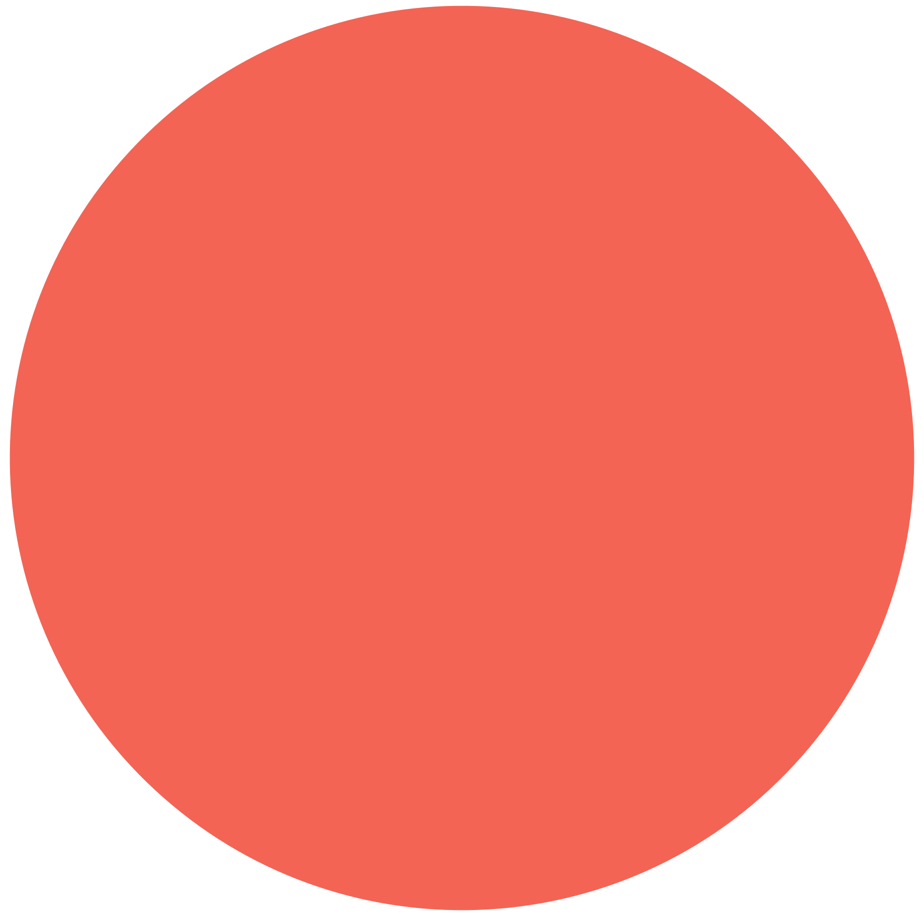


Veltmaat, Niemeyer, Schwabe (18)



An intuitive understanding of the granule structure:

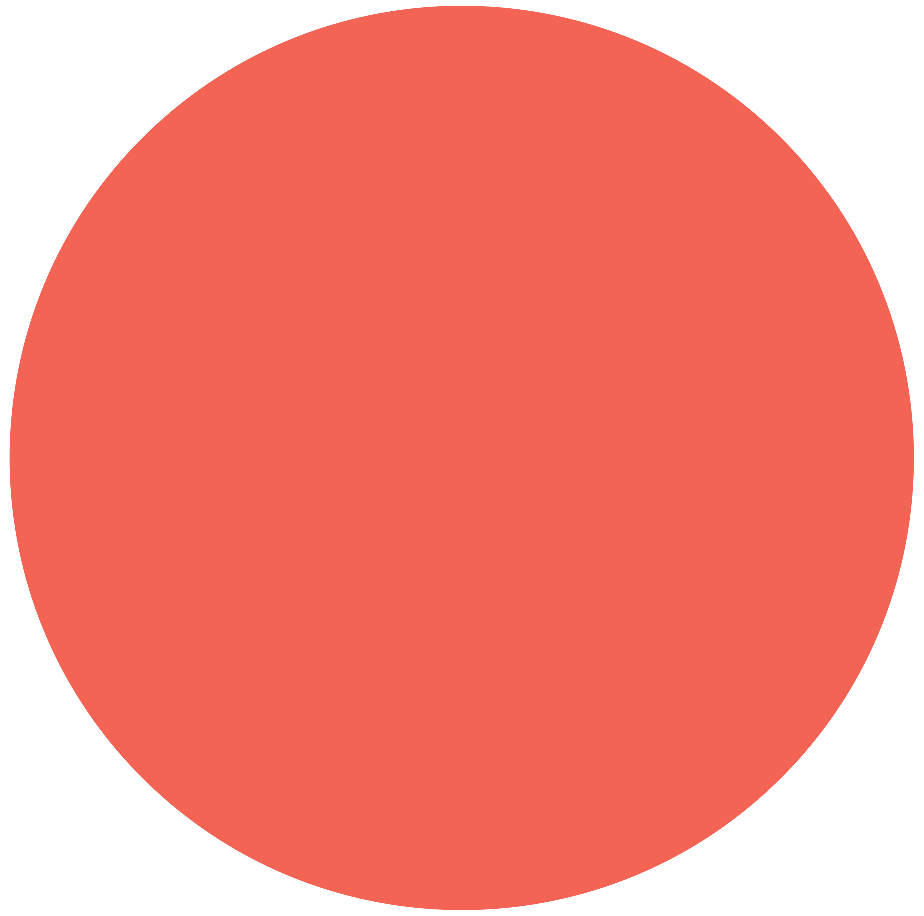
Quasiparticle



$$\ell \sim \lambda = \frac{1}{mv}$$

$$m_{\rm eff} \sim \rho_{\rm DM} \ell^3$$

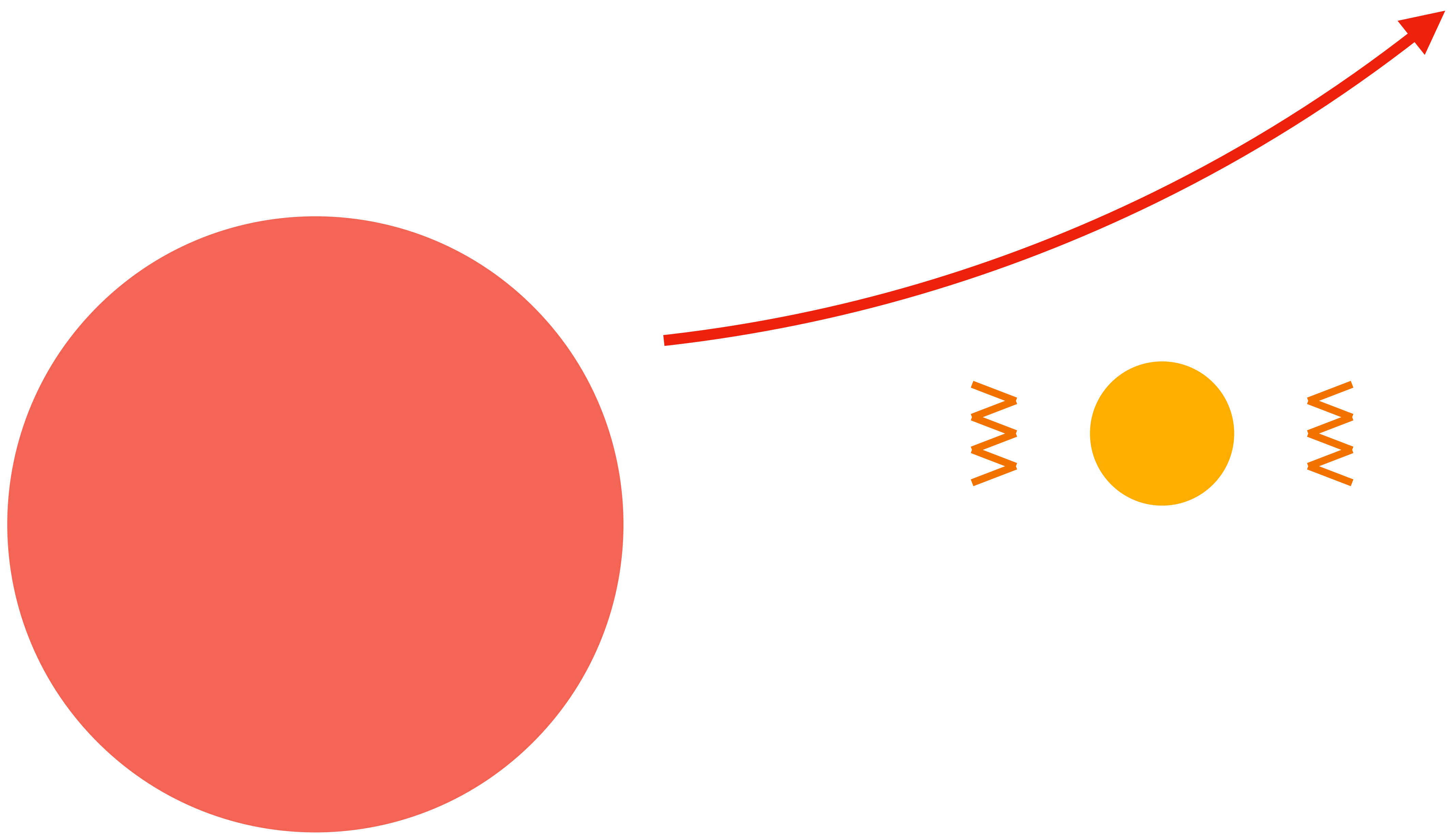
the size and mass of them could be astronomical



$$\ell \sim \lambda = \frac{1}{mv} \sim 10\,\mathrm{AU} \times \left(\frac{10^{-16}\,\mathrm{eV}}{m}\right)$$

$$m_{\mathrm{eff}} \sim \rho_{\mathrm{DM}}\ell^3 \sim 10^{15}\,\mathrm{kg} \times \left(\frac{10^{-16}\,\mathrm{eV}}{m}\right)^3$$

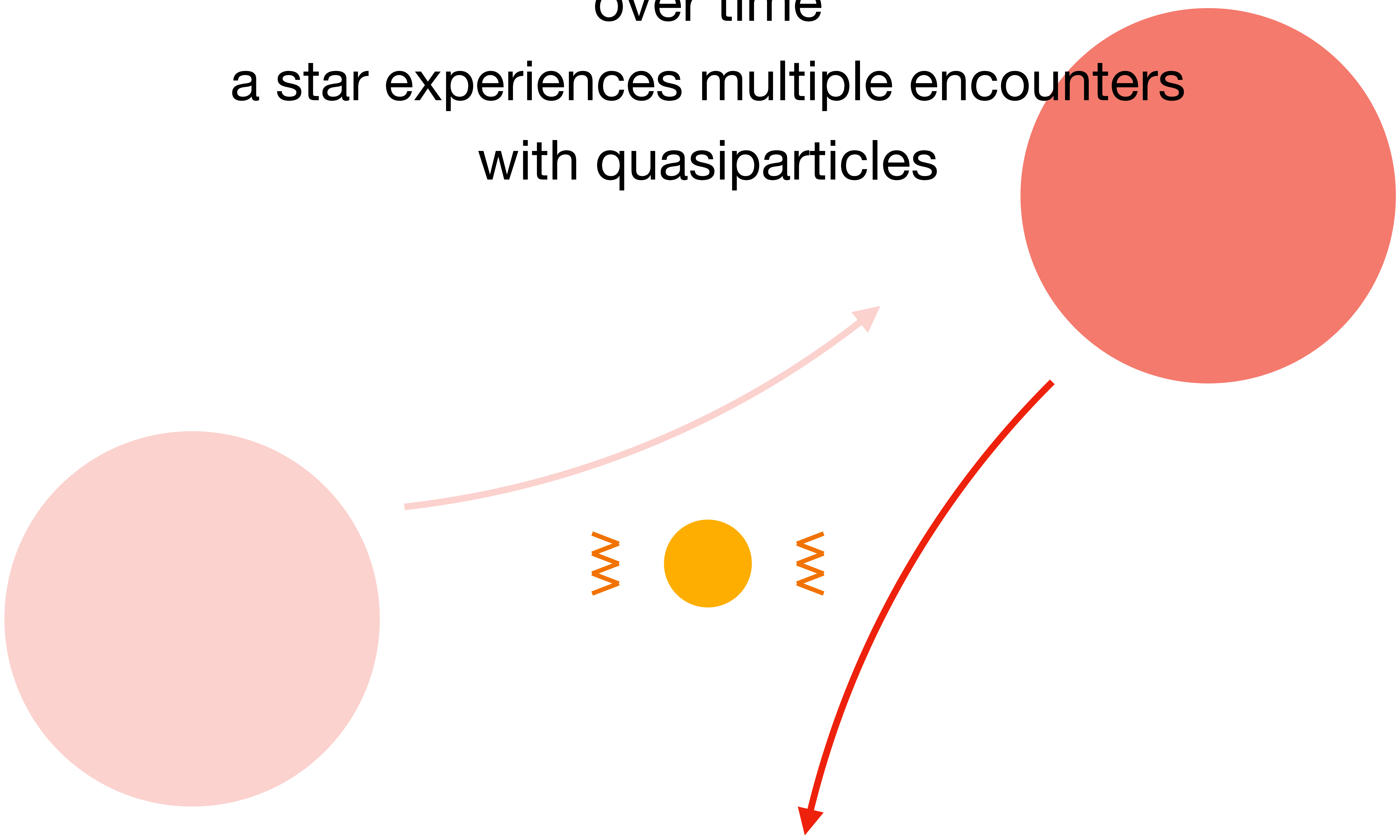
being that massive
it may engage in interaction with stars
and significantly perturb the motion of them



over time

a star experiences multiple encounters

with quasiparticles



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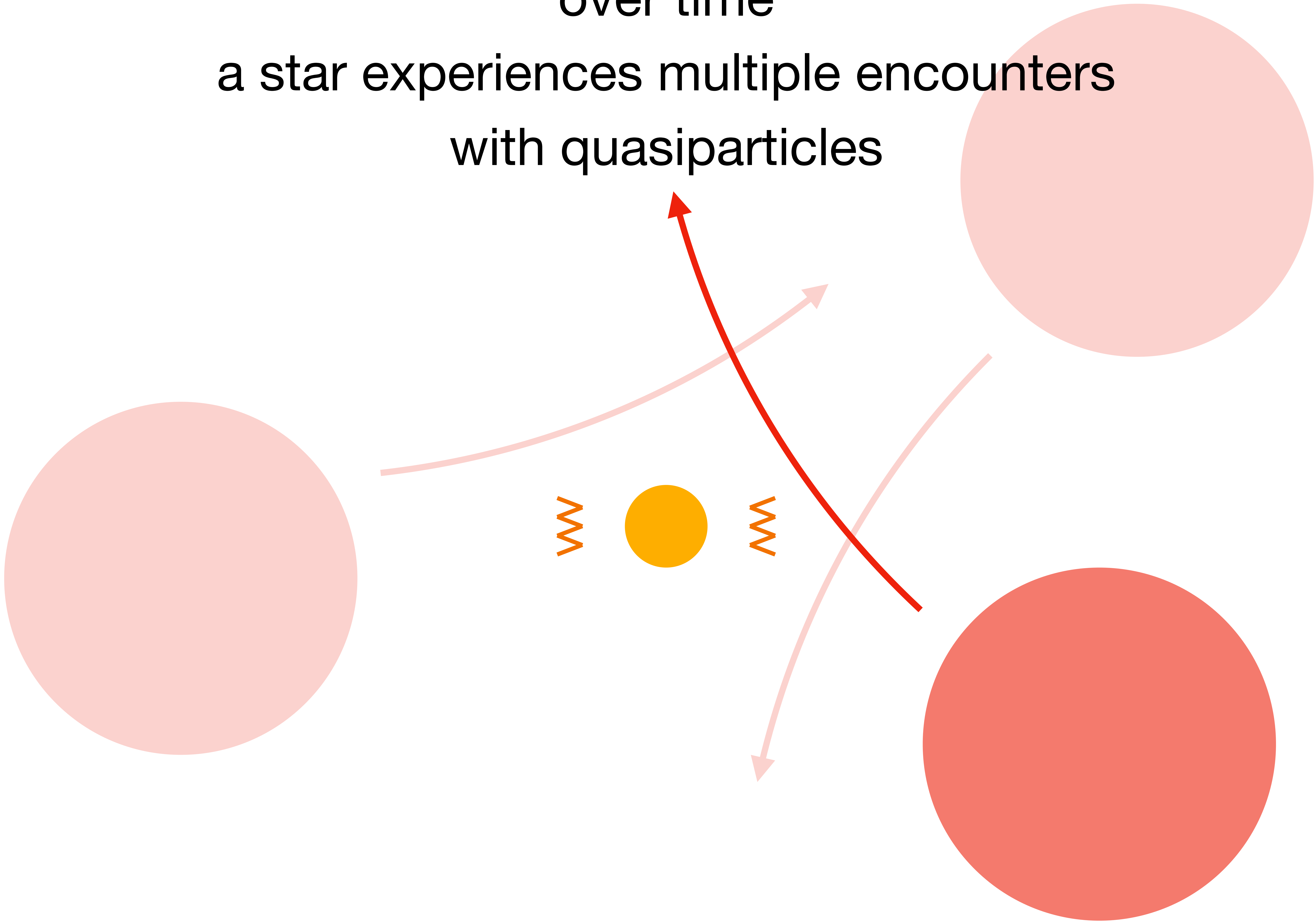
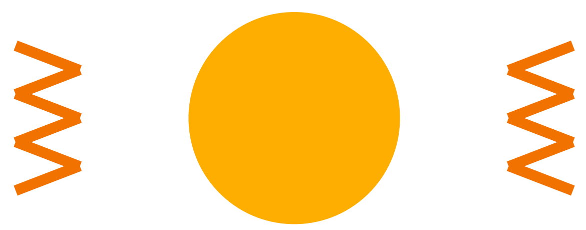


$$a = \frac{Gm_{\text{eff}}}{\lambda^2}$$

$$\tau = \frac{1}{mv^2}$$

over time
a star experiences multiple encounters
with quasiparticles



Number of encounters over T

$$\langle v^2 \rangle = (a\tau)^2 \left(\frac{T}{\tau}\right)$$

typical velocity kick over one coherence time scale so what?

so what?

quasiparticles *bombards*
normal matters, leaving *distinctive stochastic signals*
in *gravitational wave detectors*

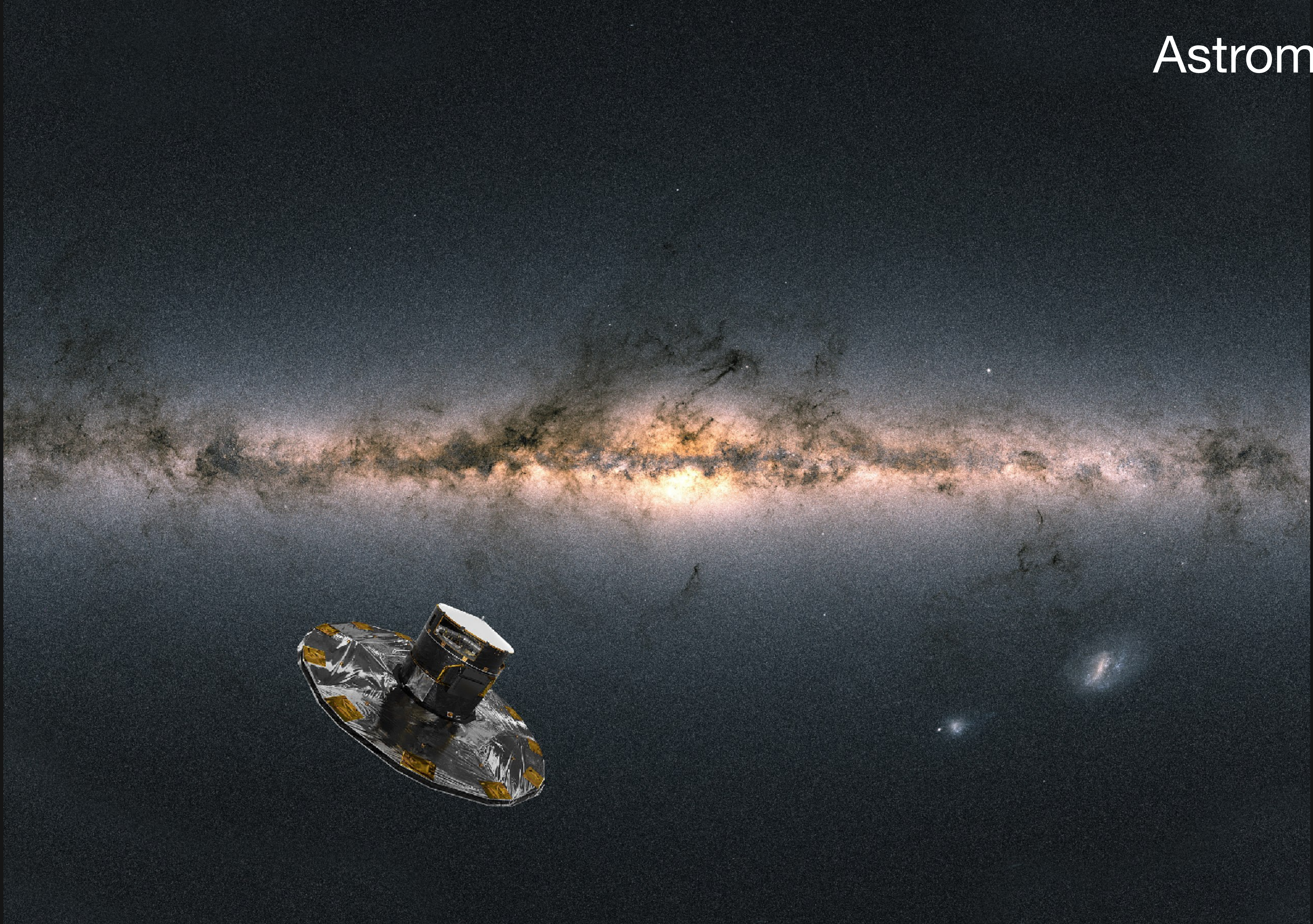
so what?

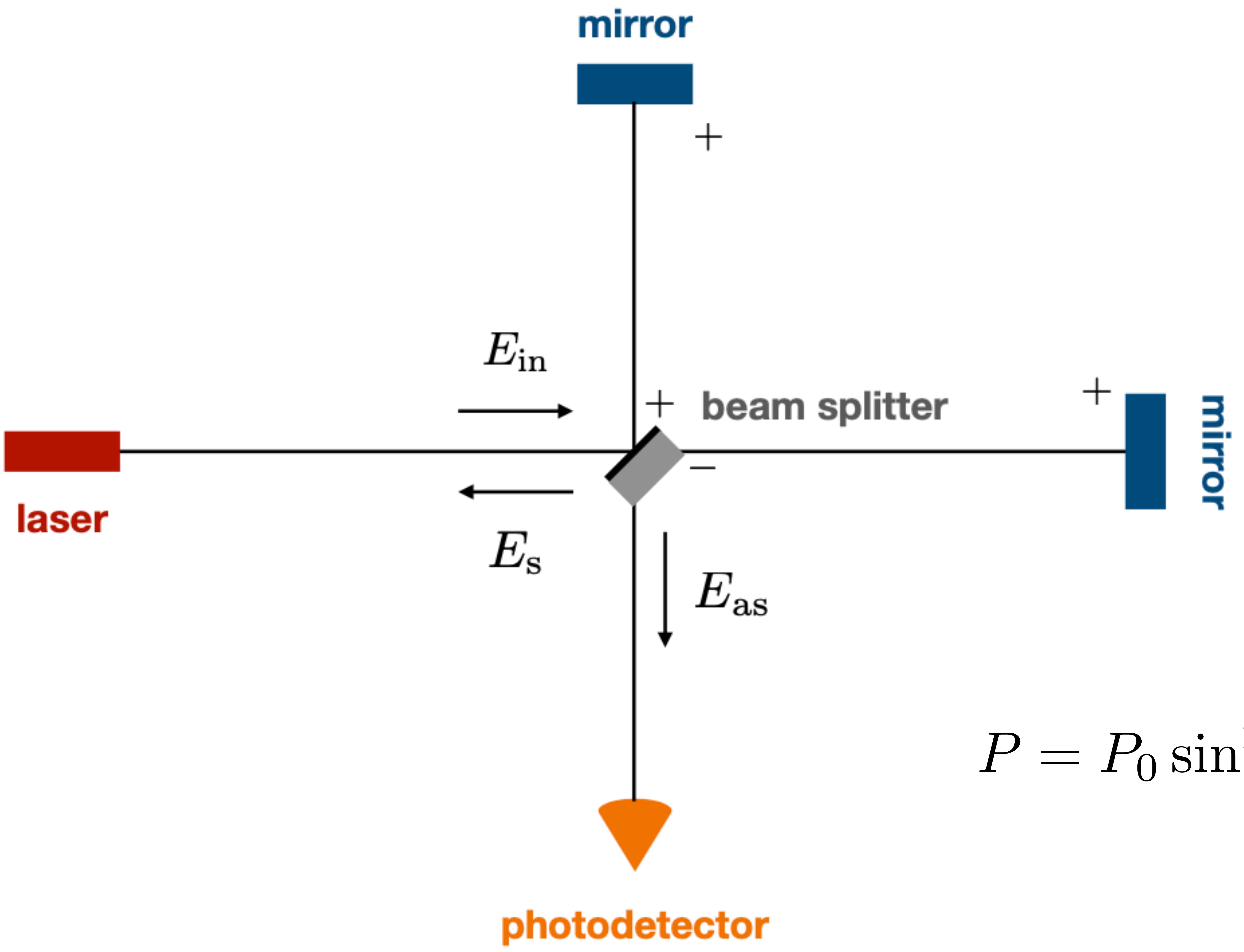
quasiparticles *bombards*
normal matters, leaving *distinctive stochastic signals*
in *gravitational wave detectors*

Interferometry

Pulsar Timing Array

Astrometry

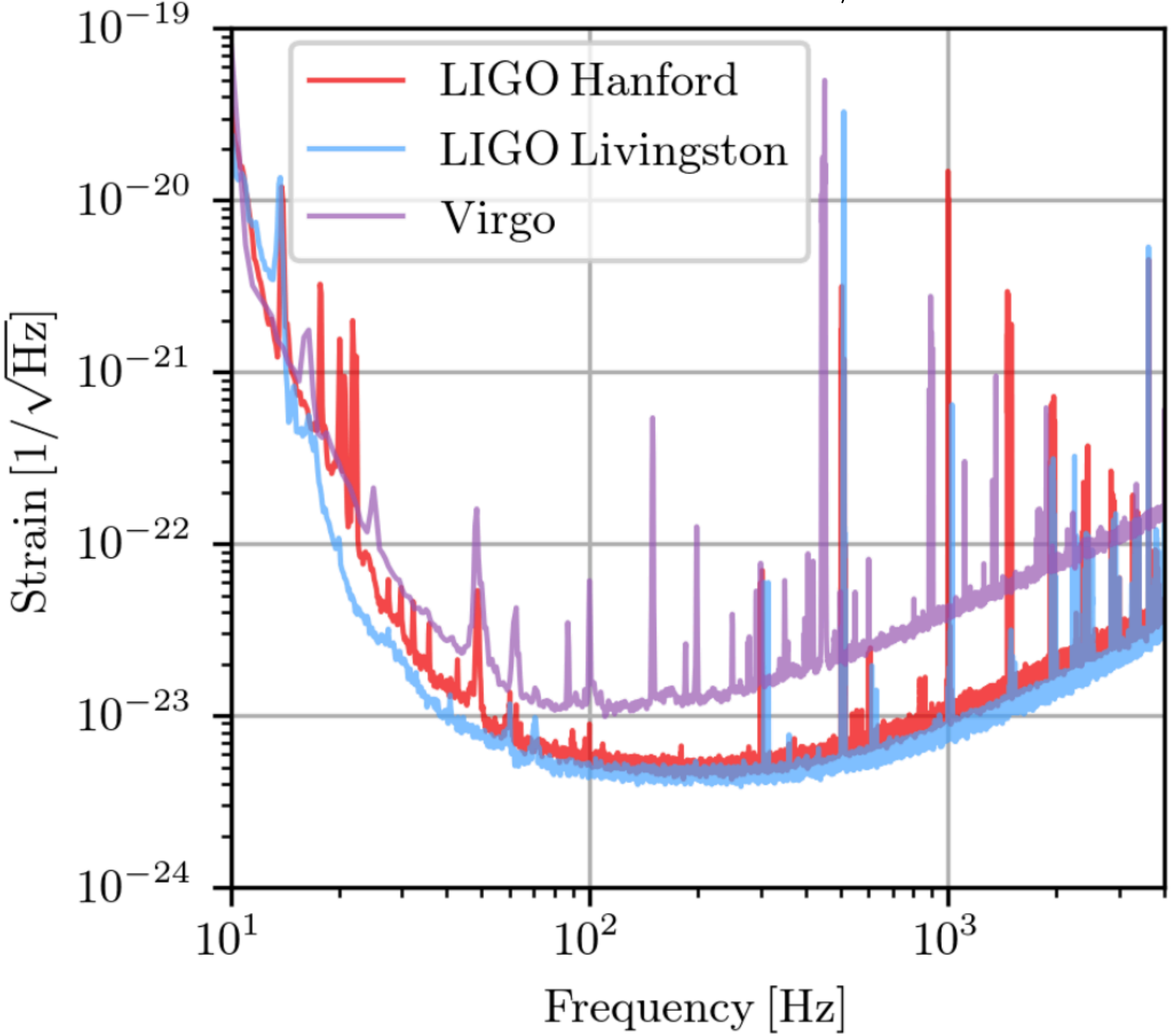


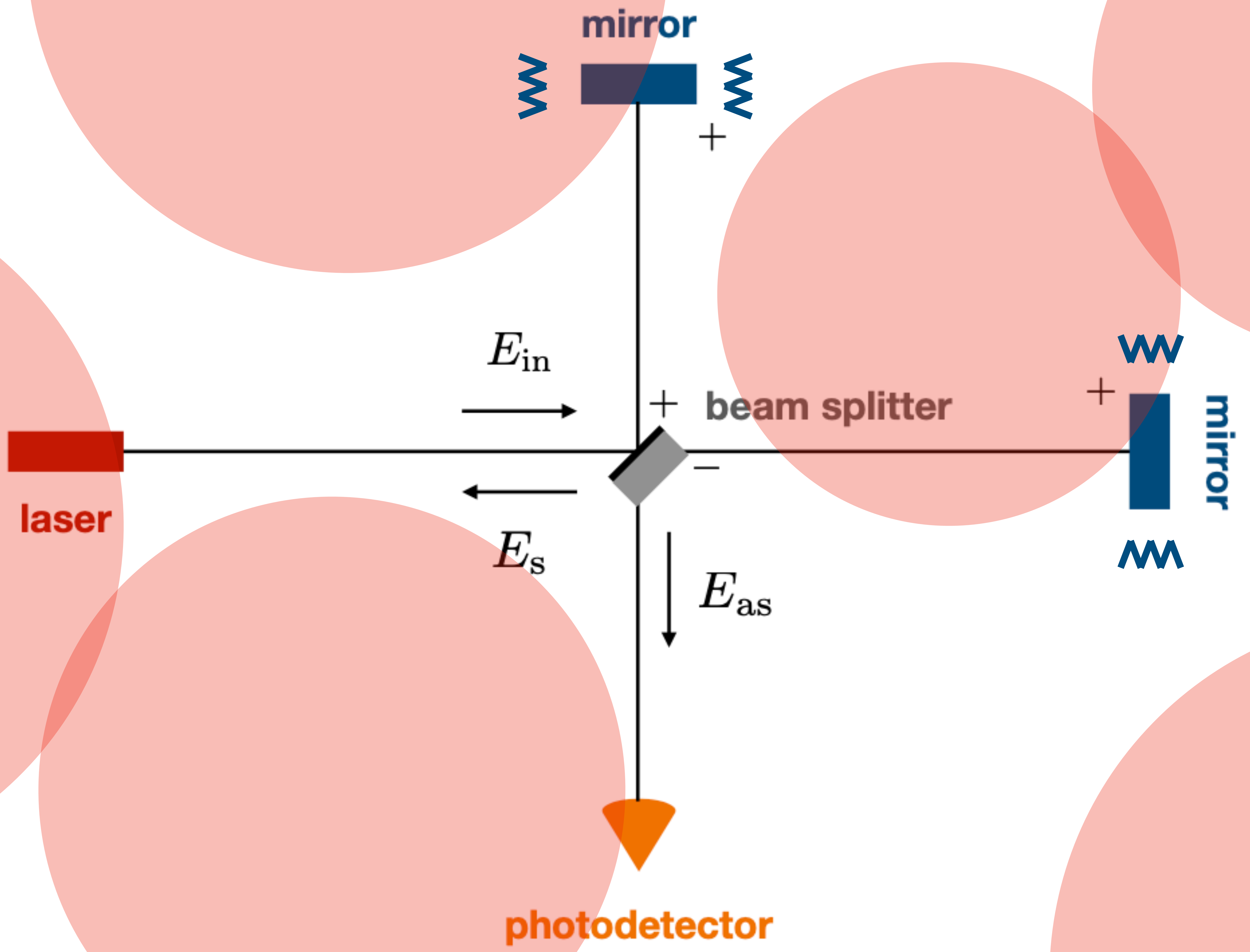


$$P = P_0 \sin^2(k\Delta L)$$

$$S_n^{1/2}(f) \sim S_{\Delta L/L}^{1/2}(f)$$

$$\langle x^2 \rangle = \int df\, S_x(f)$$





mirror
+
E_{in}
+
beam splitter
+
mirror
laser
−
E_{s}
E_{as}
photodetector

can we actually measure
ULDM signals with GW interferometers?

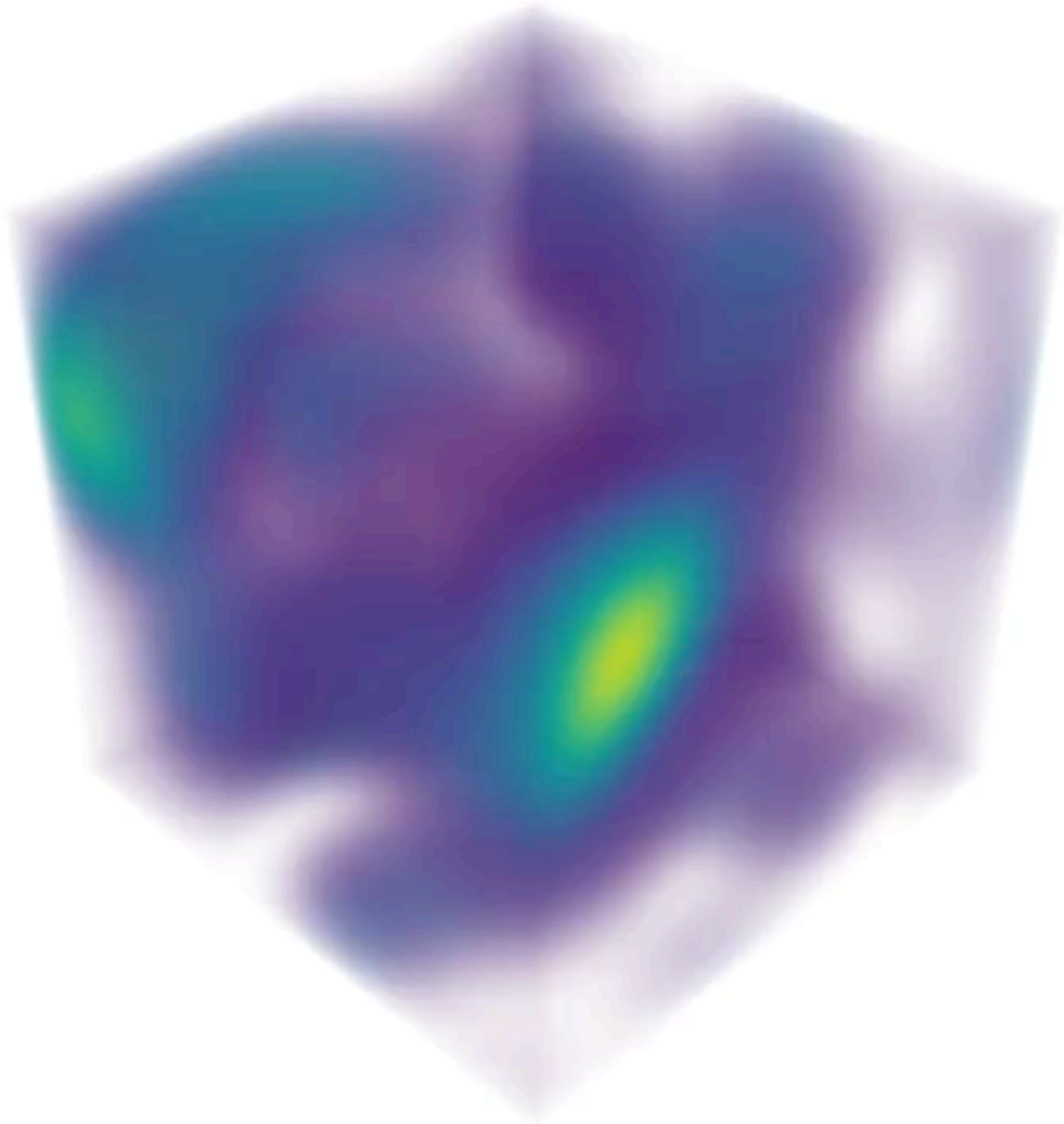
$$\ddot{x} = -\nabla\Phi$$

$$\nabla^2\Phi = 4\pi G\rho$$

what is reflected in *detector observables*

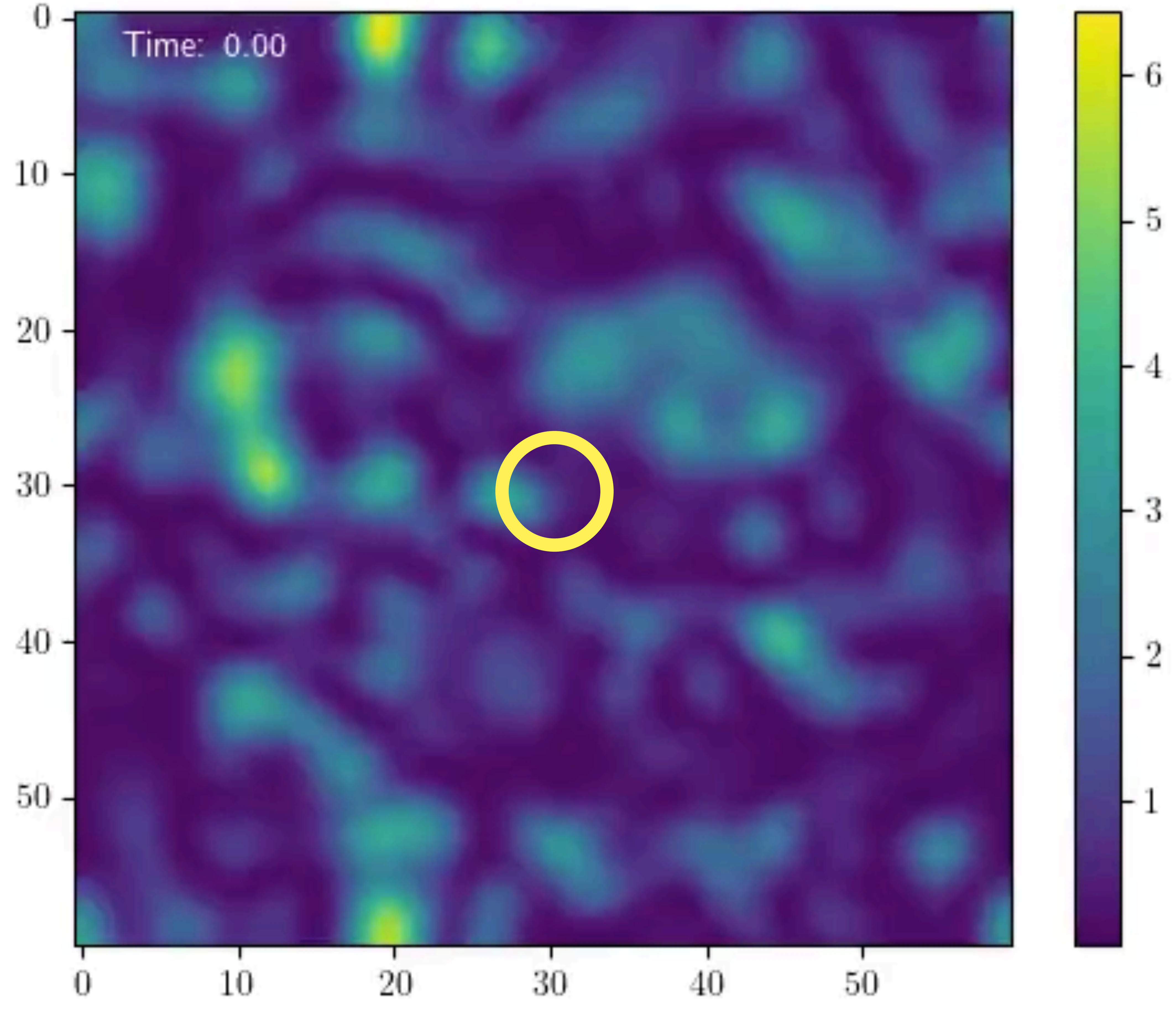
is the *statistical properties of*

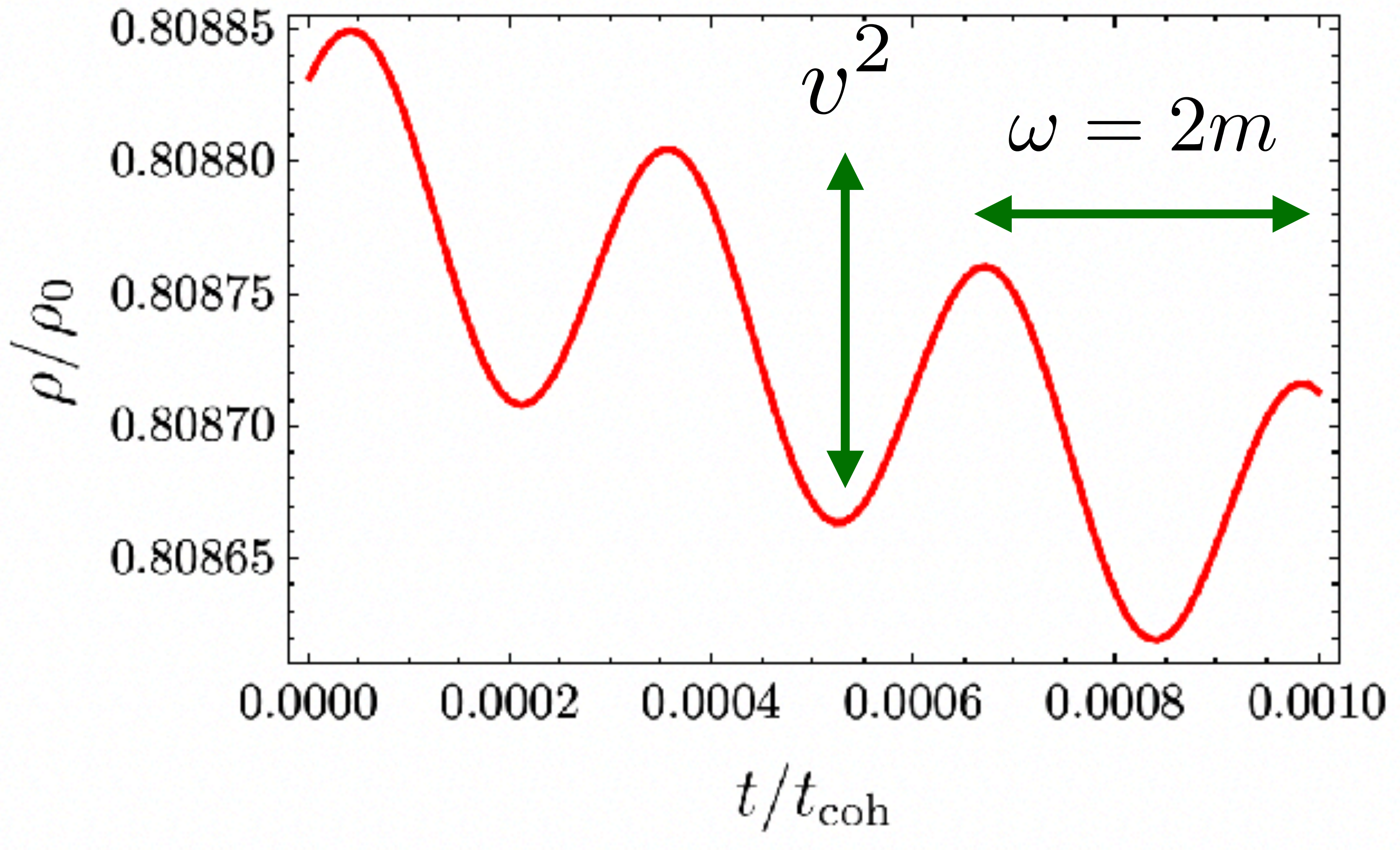
density fluctuations of ULDM

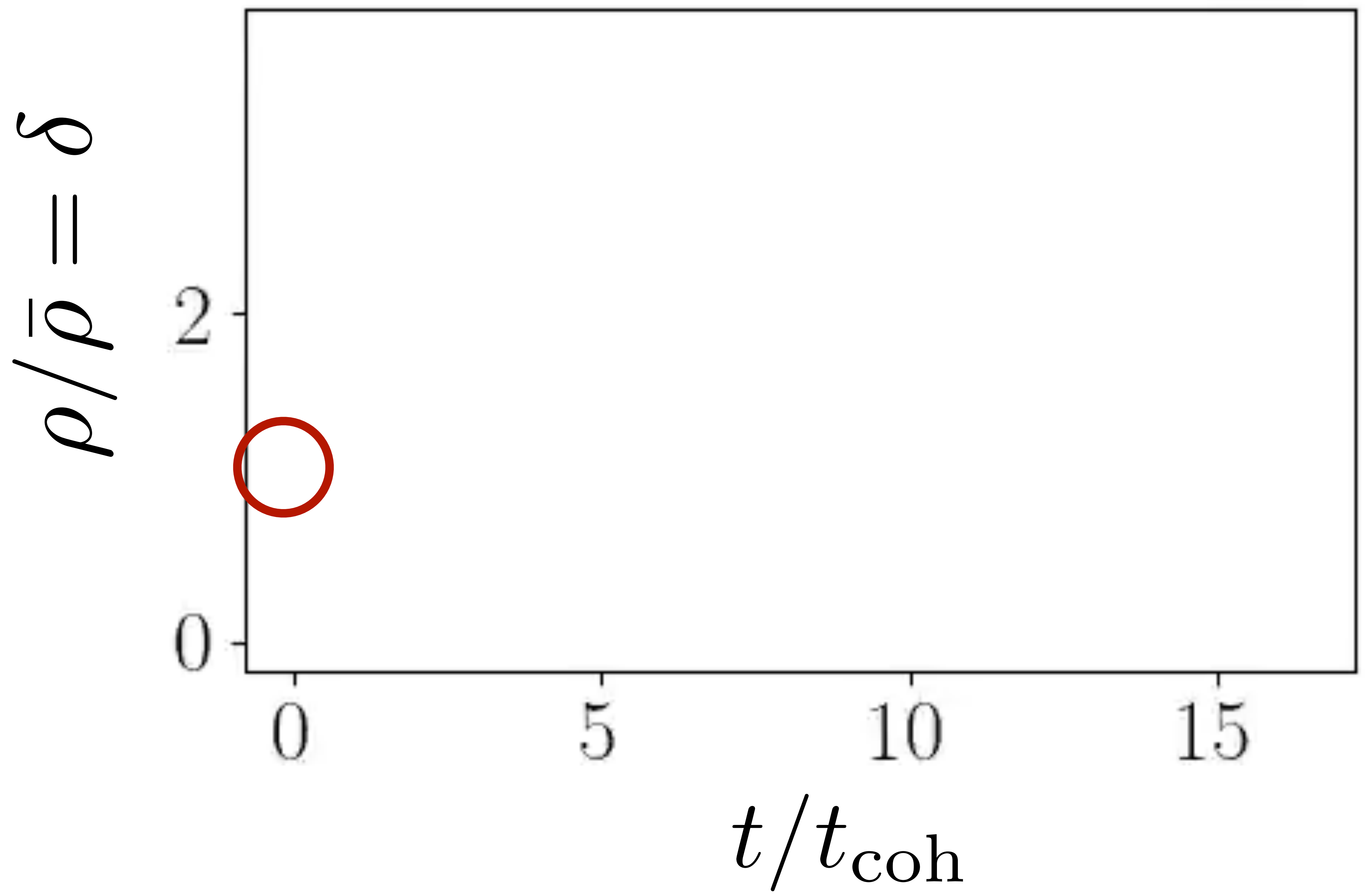


the density-density correlator at the same position is

$$\langle \delta(x)\delta(x) \rangle = \int \frac{d\omega}{2\pi} S_\delta(\omega)$$







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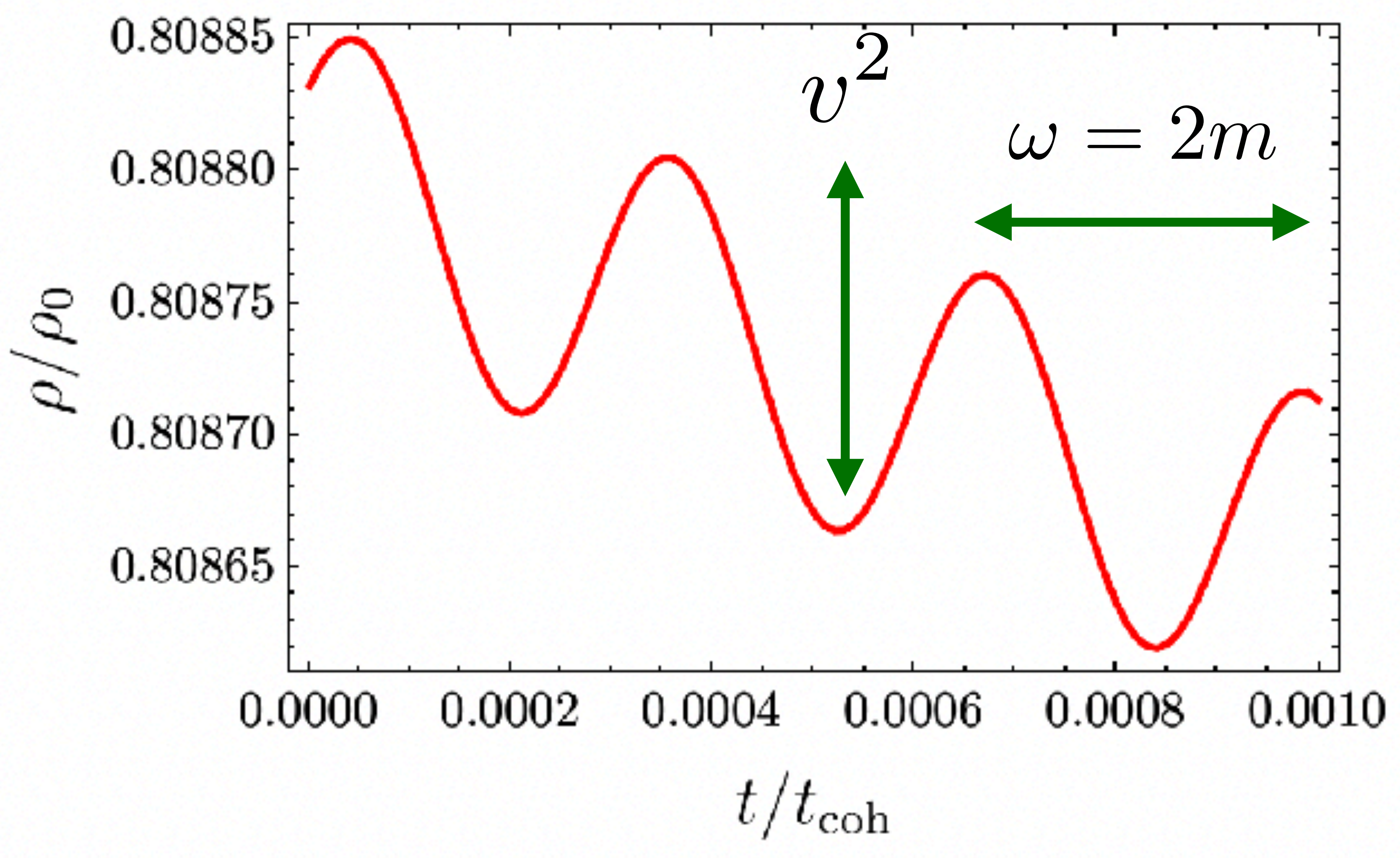
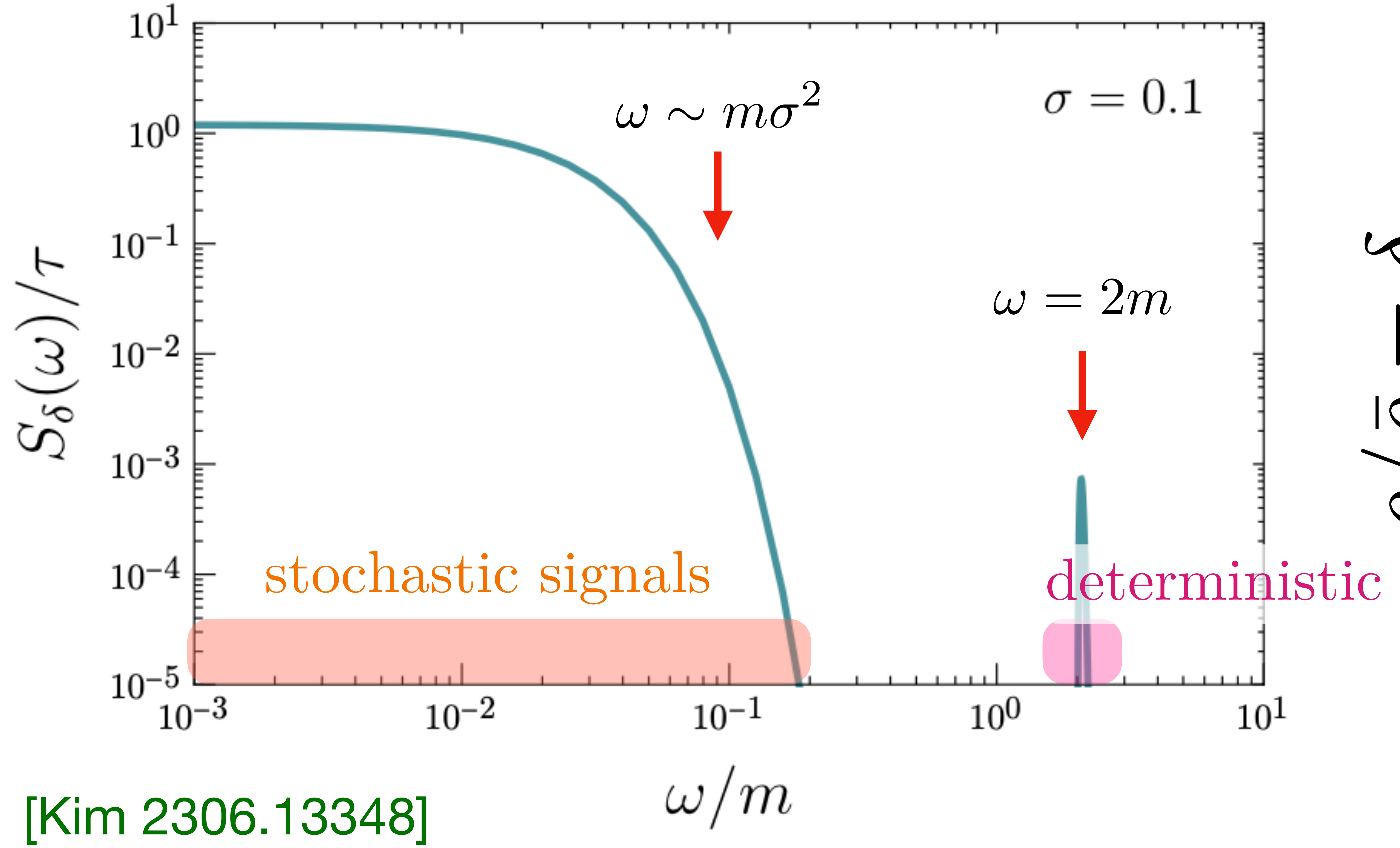
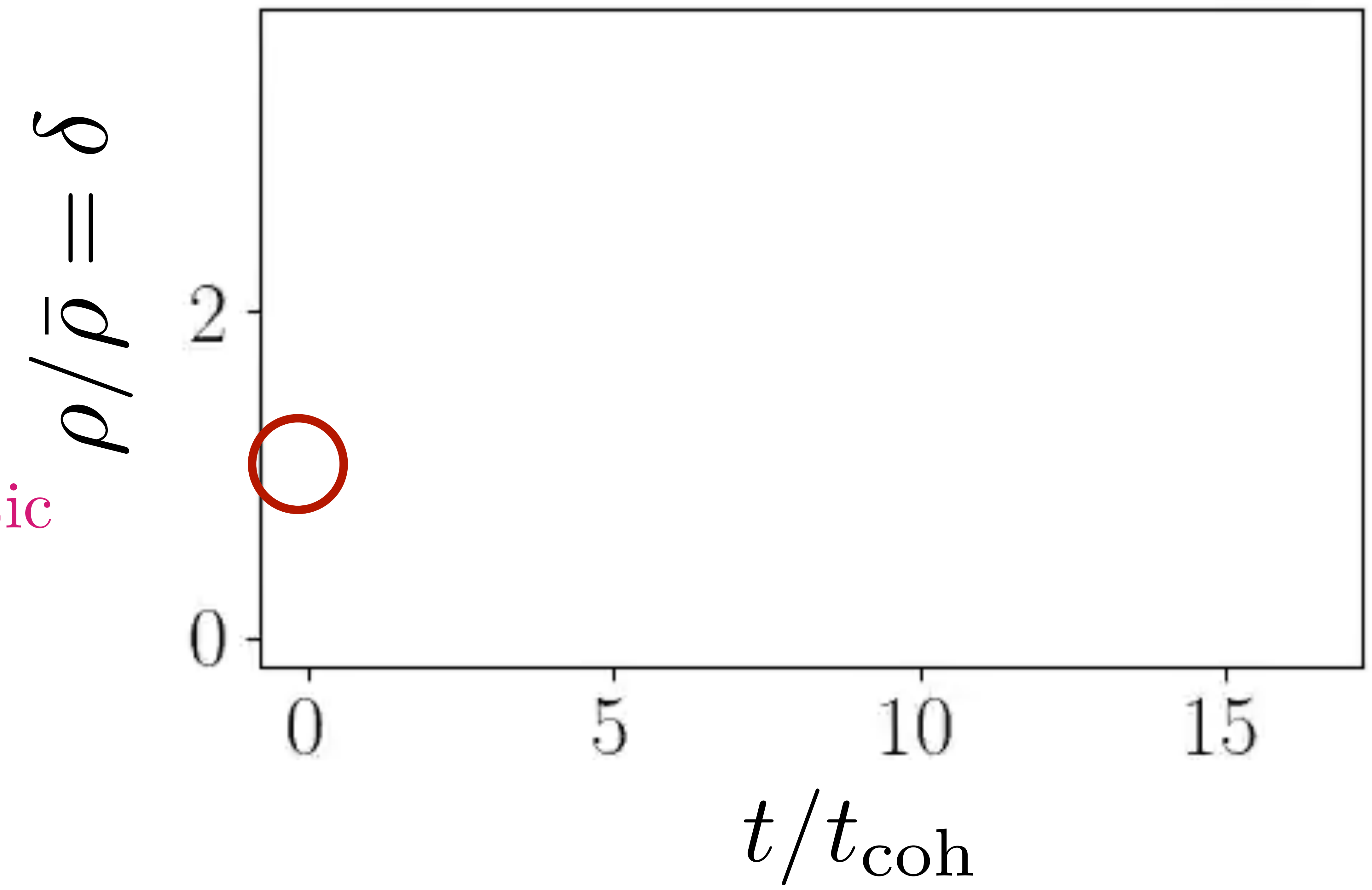


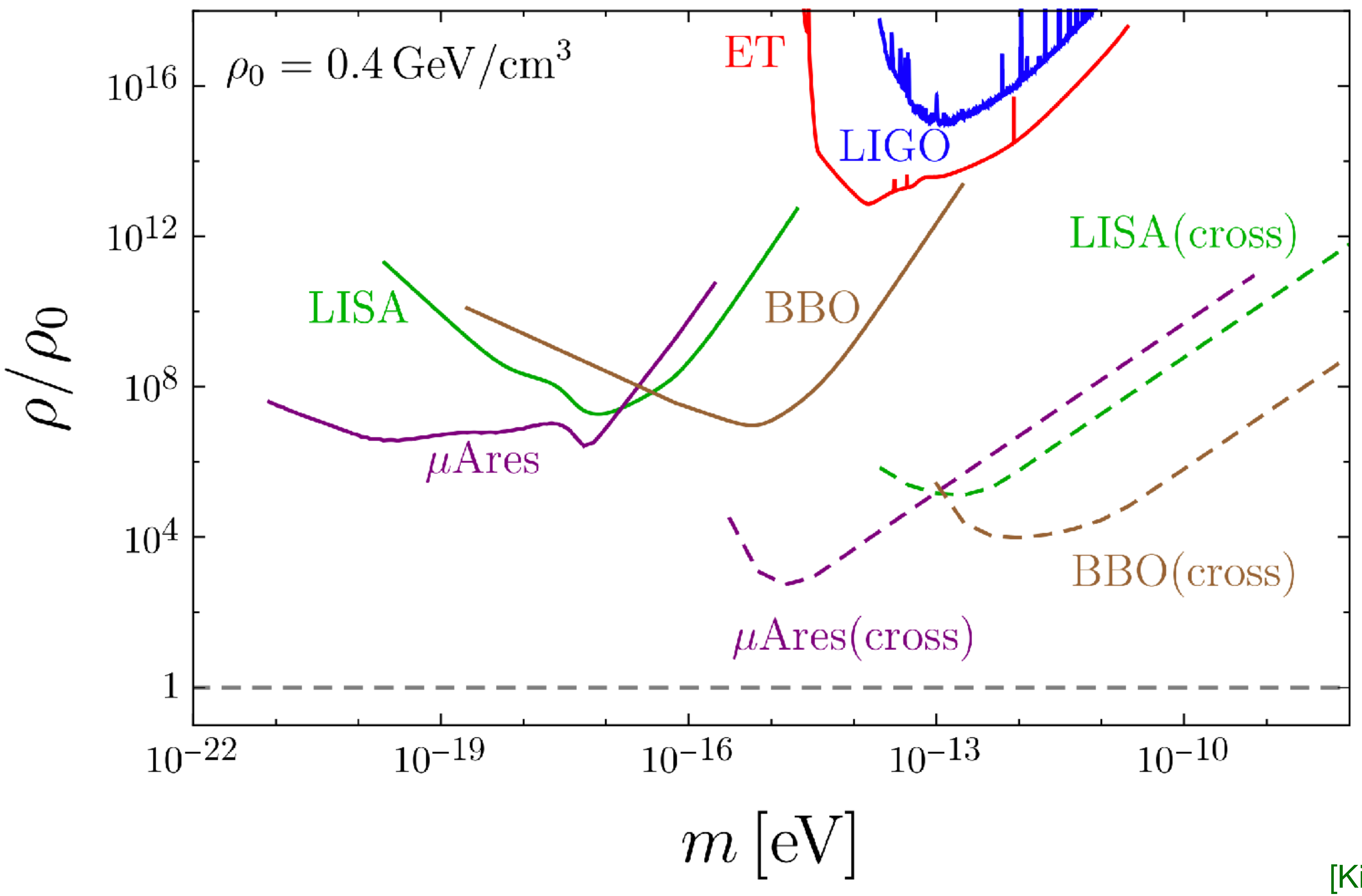
$$S_\delta(\omega) = \tau[\sigma^4 A_\delta(\omega) + B_\delta(\omega)]$$





[Kim 2306.13348]

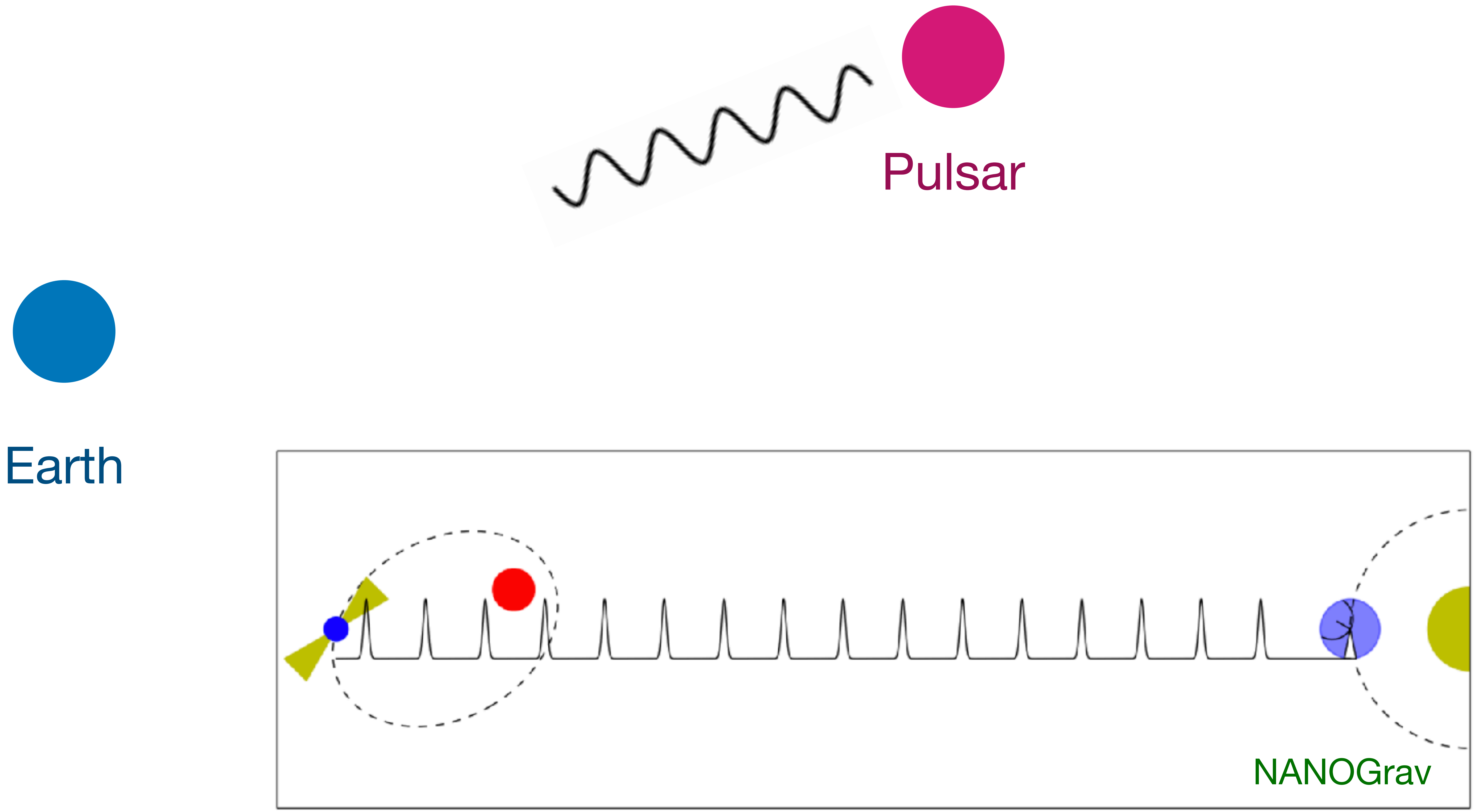
[Kim, Lenoci, Perez, Ratzinger, 2307.14962]

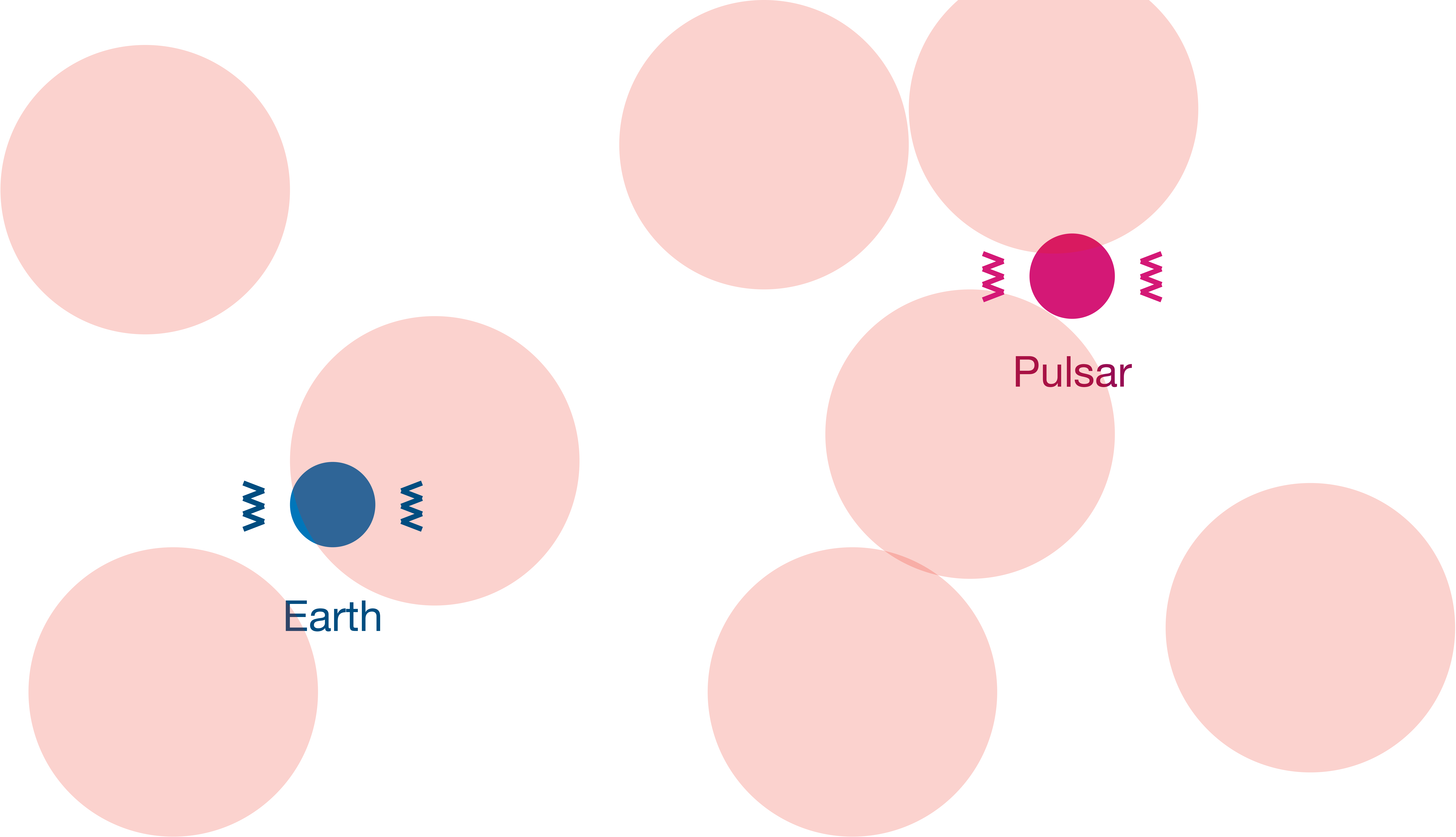


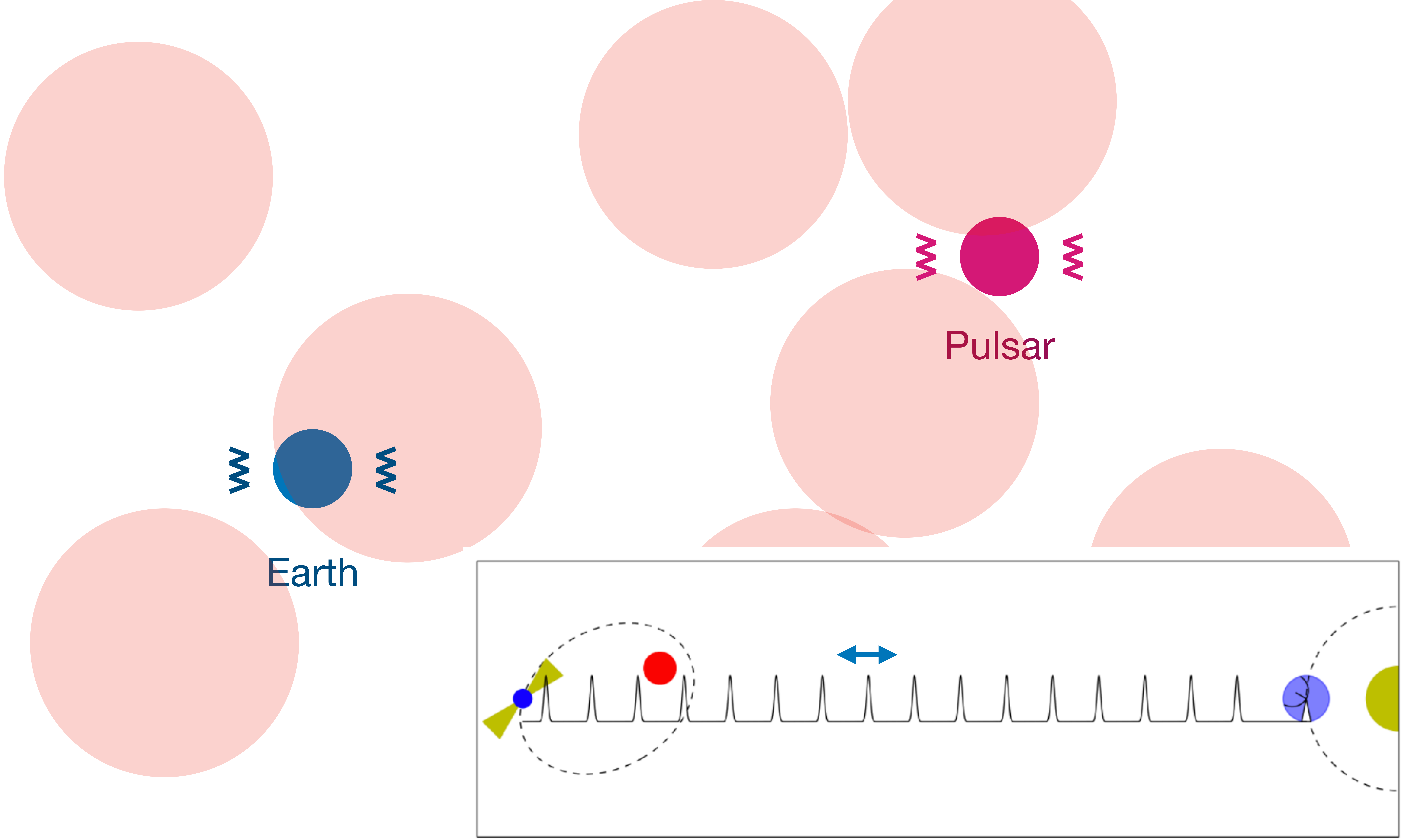
[Kim 2306.13348]

another example:

Pulsar Timing Array


Pulsar
Earth
NANOGrav


Pulsar
Earth


Pulsar
Earth

PTA is a natural place to look for ULDM signals:

I. long baseline (Earth-Pulsar distance ~ kpc)

II. multiple pulsars are monitored; correlation can be used ultralight dark matter signal is characterised by

spectrum and ***correlation***

$$\langle \delta t_a \delta t_b \rangle = \int df\, \Gamma_{ab}^{\mathrm{ULDM}} S_{\delta t}^{\mathrm{ULDM}}(f)$$

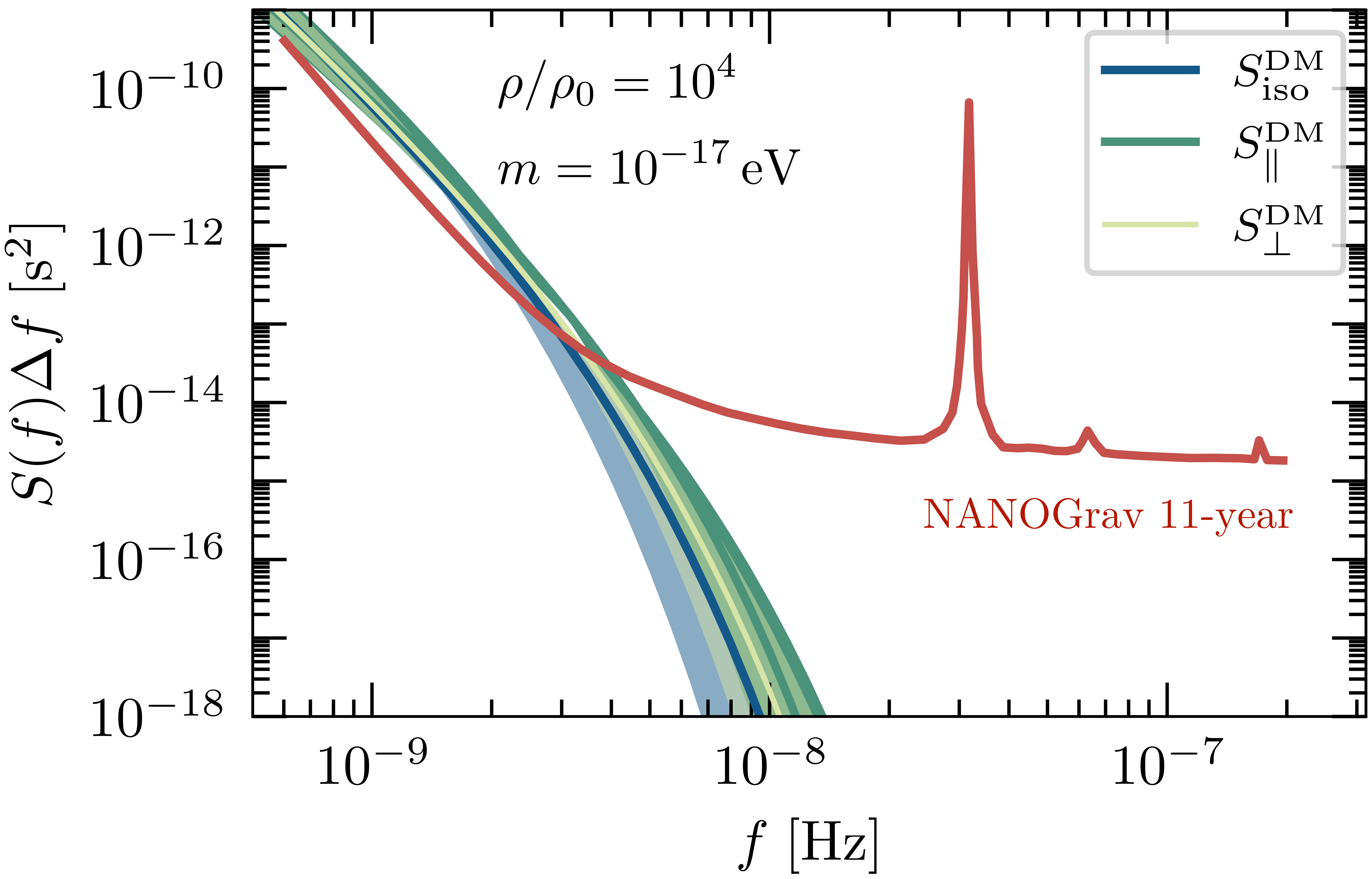
$$\langle \delta t_a \delta t_b \rangle = \int df \, \Gamma_{ab}^{\rm ULDM} S_{\delta t}^{\rm ULDM}(f)$$

for isotropic DM distribution

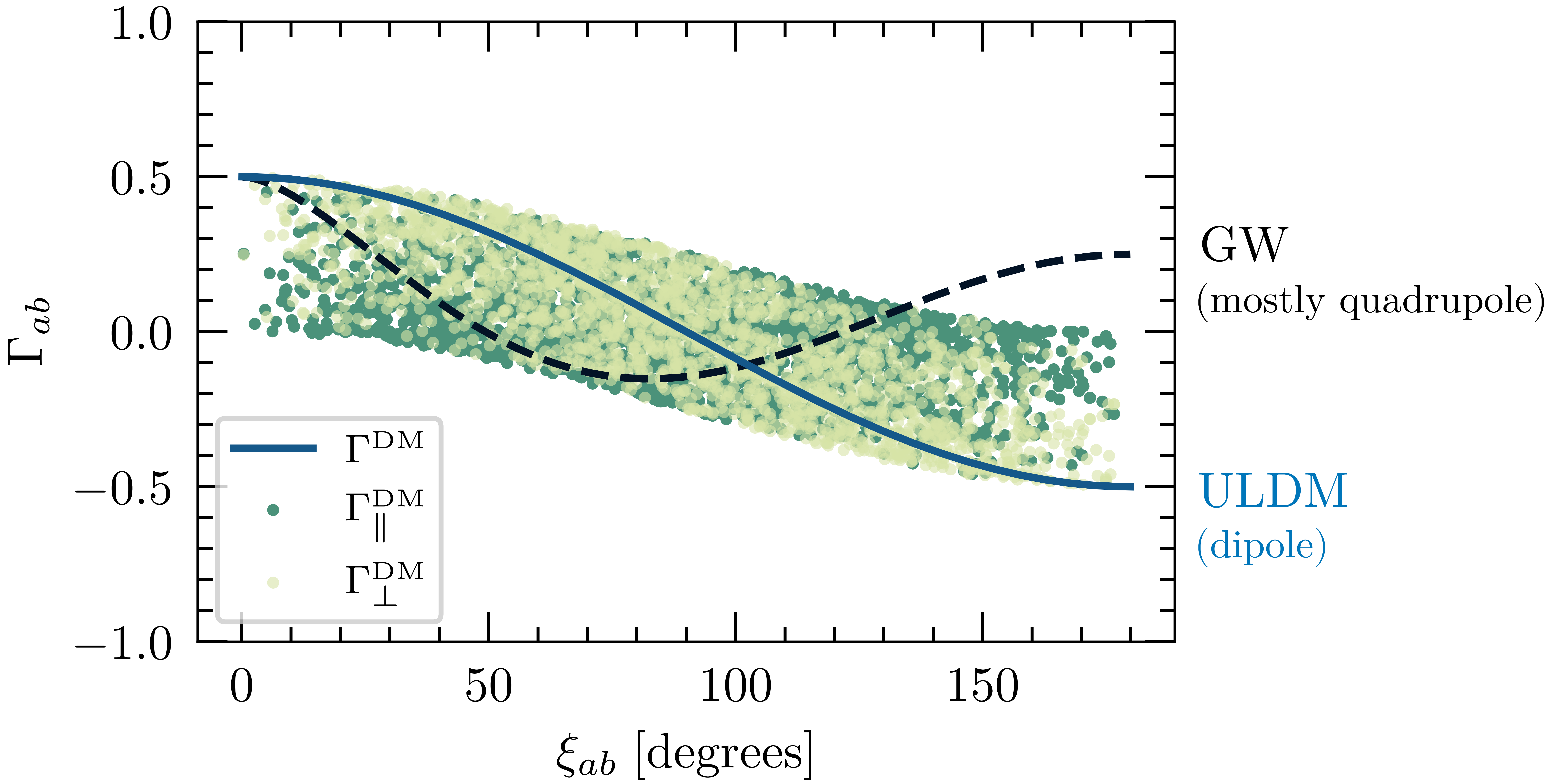
$$\Gamma_{ab} = \frac{1}{2}\left[\delta_{ab} + \hat{n}_a \cdot \hat{n}_b\right]$$

$$S_{\delta t}(f) = \frac{a^2 \tau}{(2\pi f)^4}\left[\frac{64}{3\pi} K_0(\omega/m\sigma^2)\right]$$

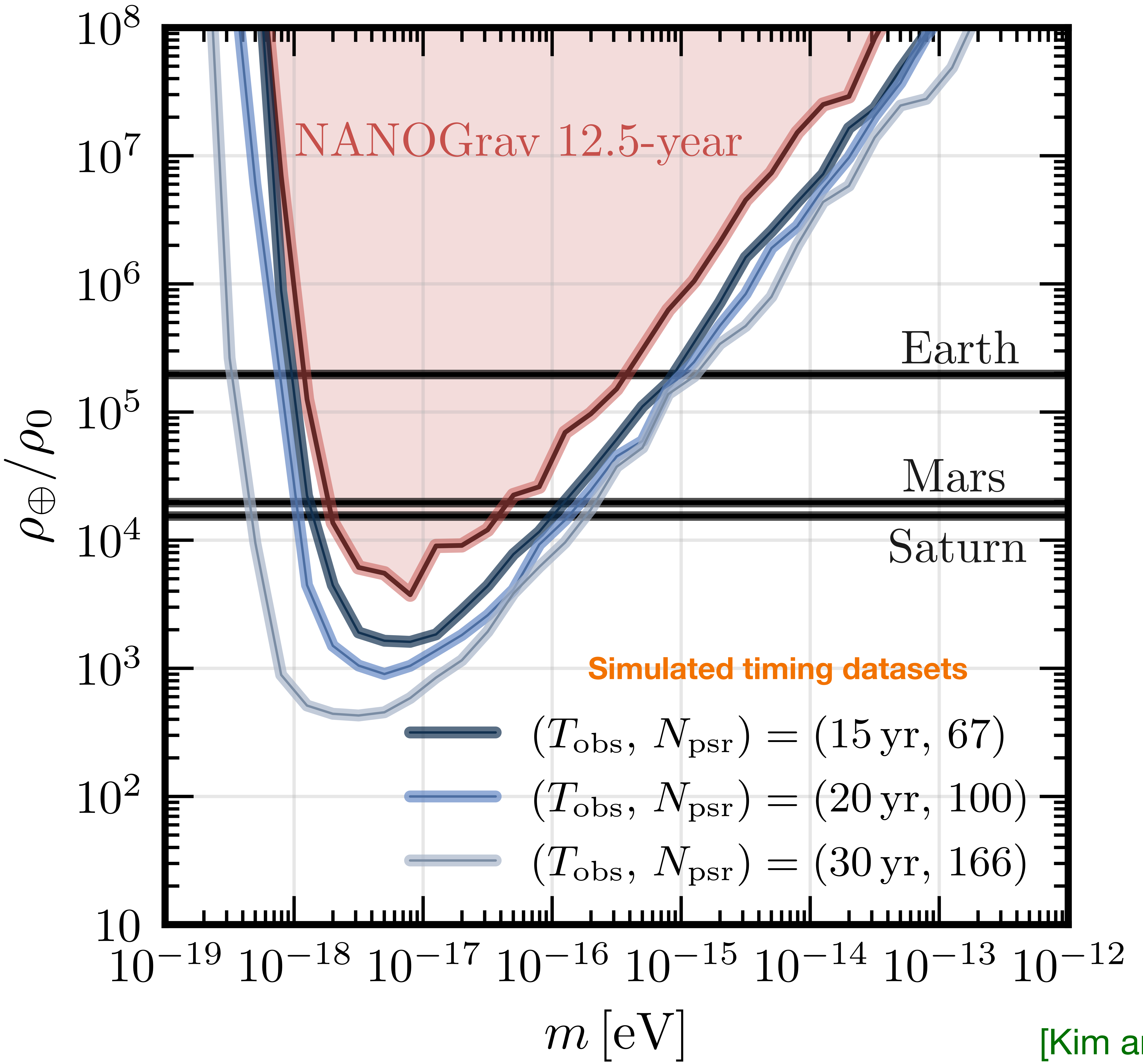
[Kim and Mitridate 2312.12225]



[Kim and Mitridate 2312.12225]



[Kim and Mitridate 2312.12225]



[Kim and Mitridate 2312.12225]

one last example:

Astrometry

astrometry involves
precision measurements of
positions / velocities of stars

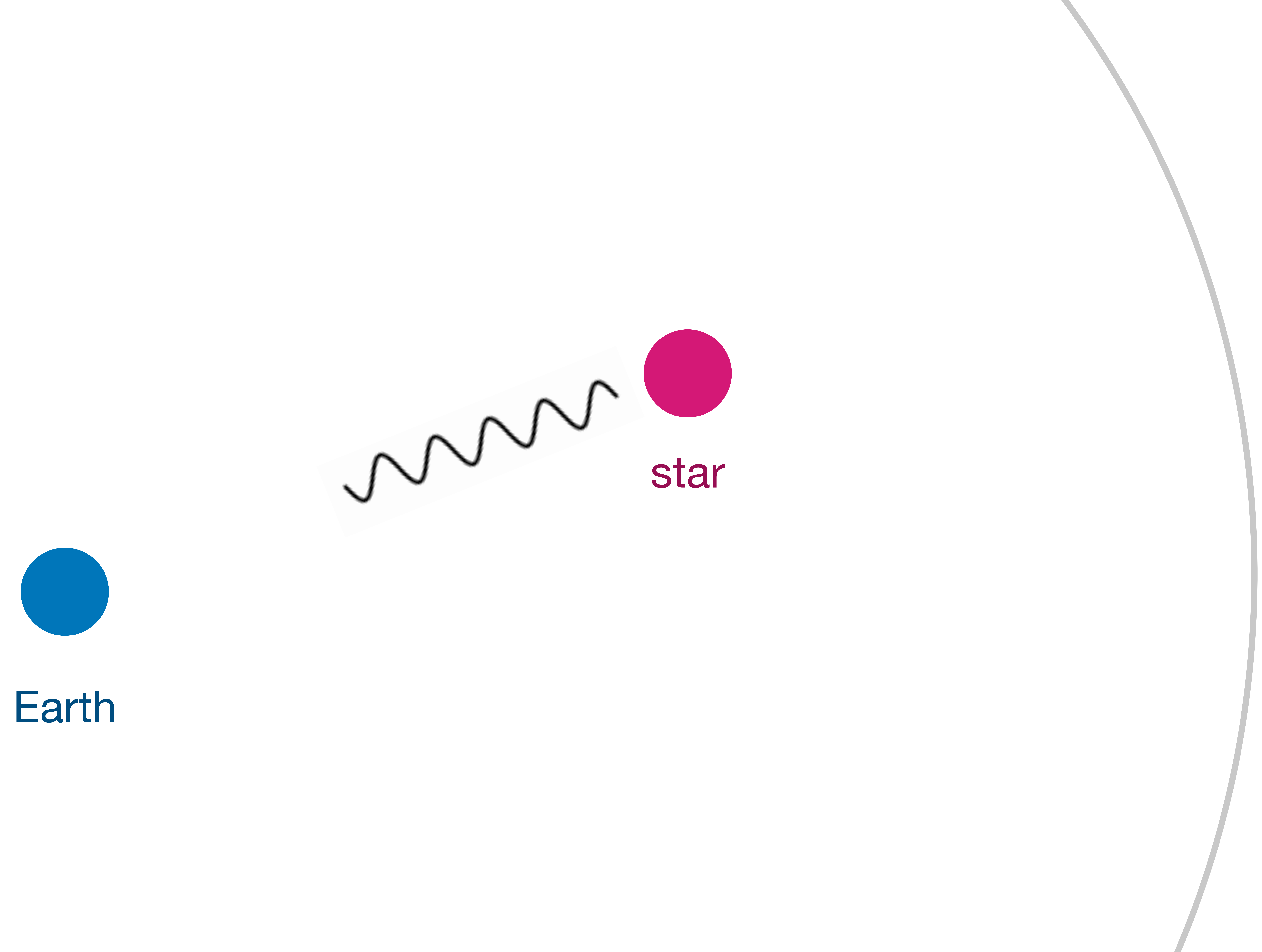
current/future astrometry missions measure

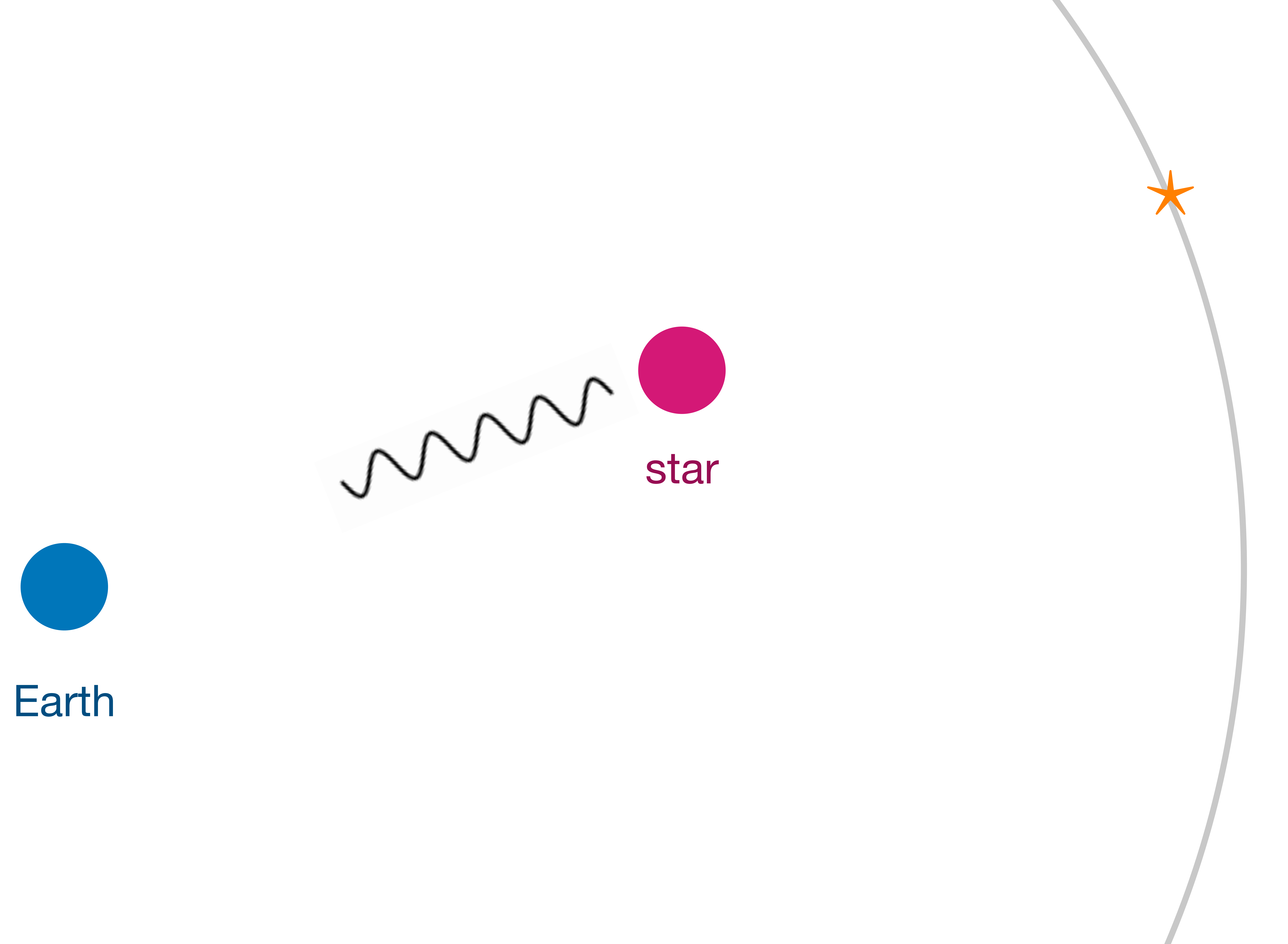
$$N_\star = 10^8 - 10^9$$

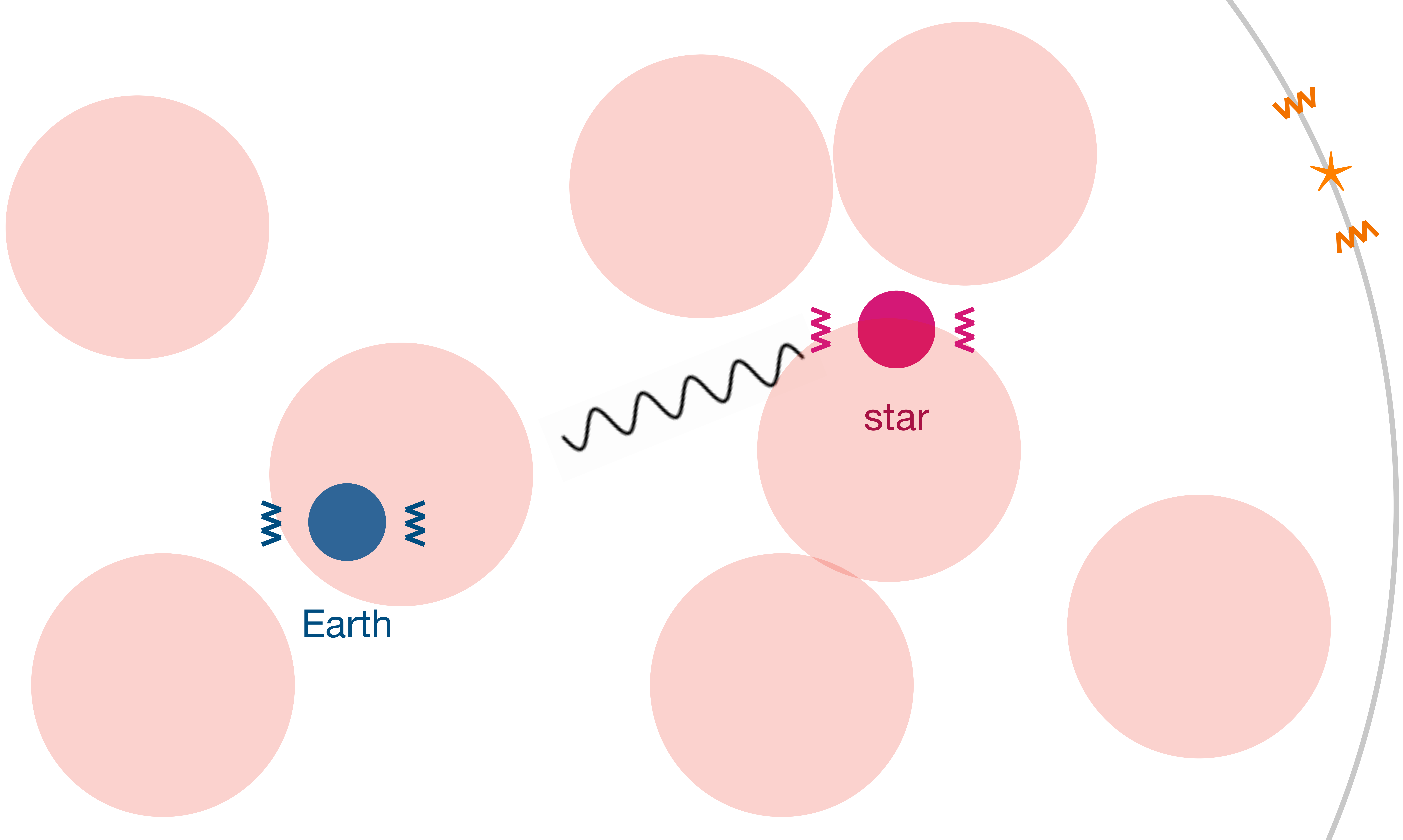
at the precision of

$$\Delta\theta \sim \mathcal{O}(10^2)\mu\mathrm{as}$$

$\mu\mathrm{as} = 5 \times 10^{-12}\,\mathrm{rad}$


star
Earth


star
Earth


star
Earth

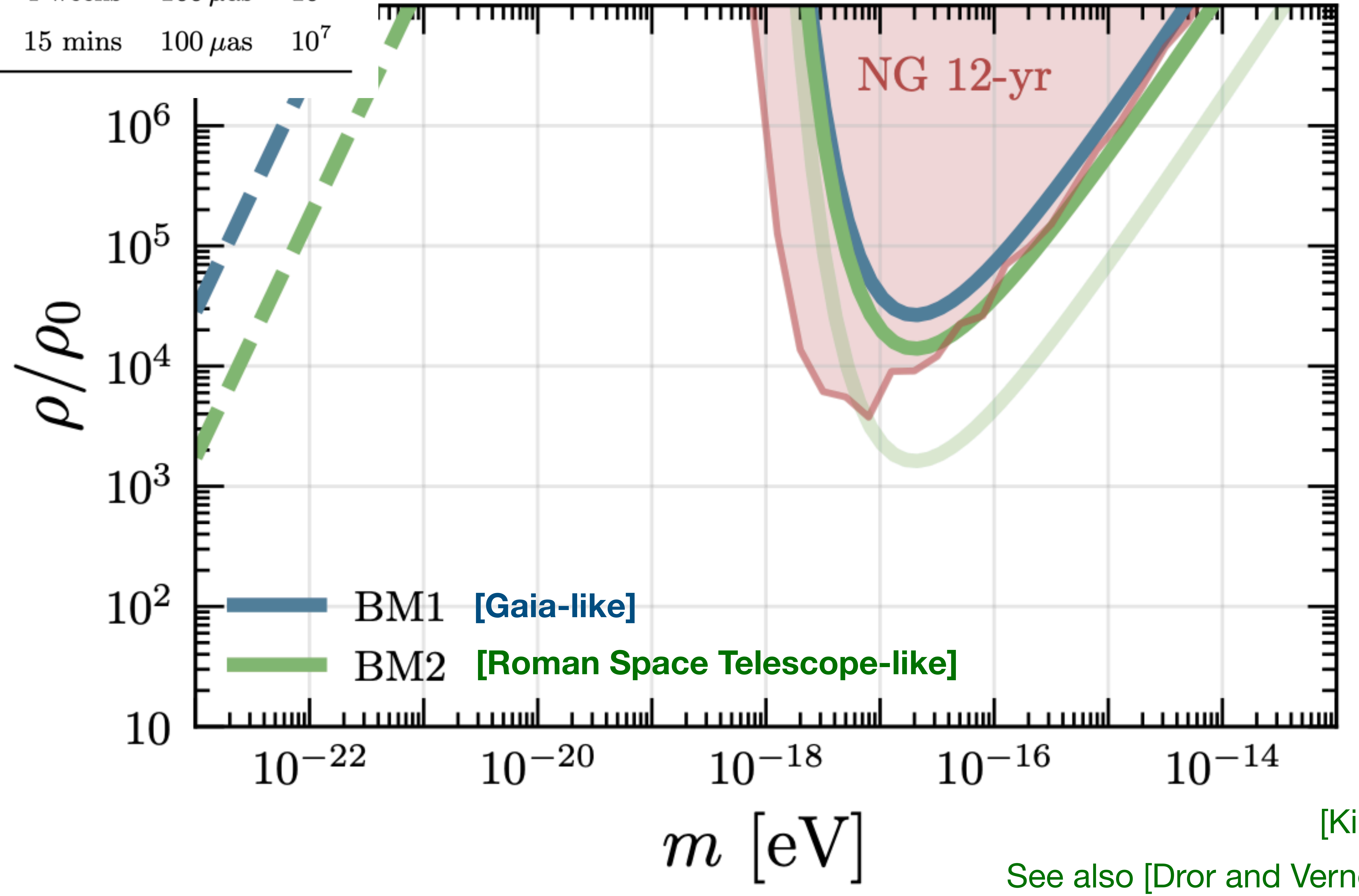
similarly

the signal is characterised by spectrum and correlation

$$\langle \delta n_a^i(t) \delta n_b^j(t') \rangle = \int df \, \Gamma_{ab}^{ij} S(f) \cos[2\pi f(t - t')]$$

[Kim 2406.03539]

	T	Δt	σ_r	$N_\star$
BM1	10-yr	4 weeks	100 μas	10^8
BM2	10-yr	15 mins	100 μas	10^7



[Kim 2406.03539]

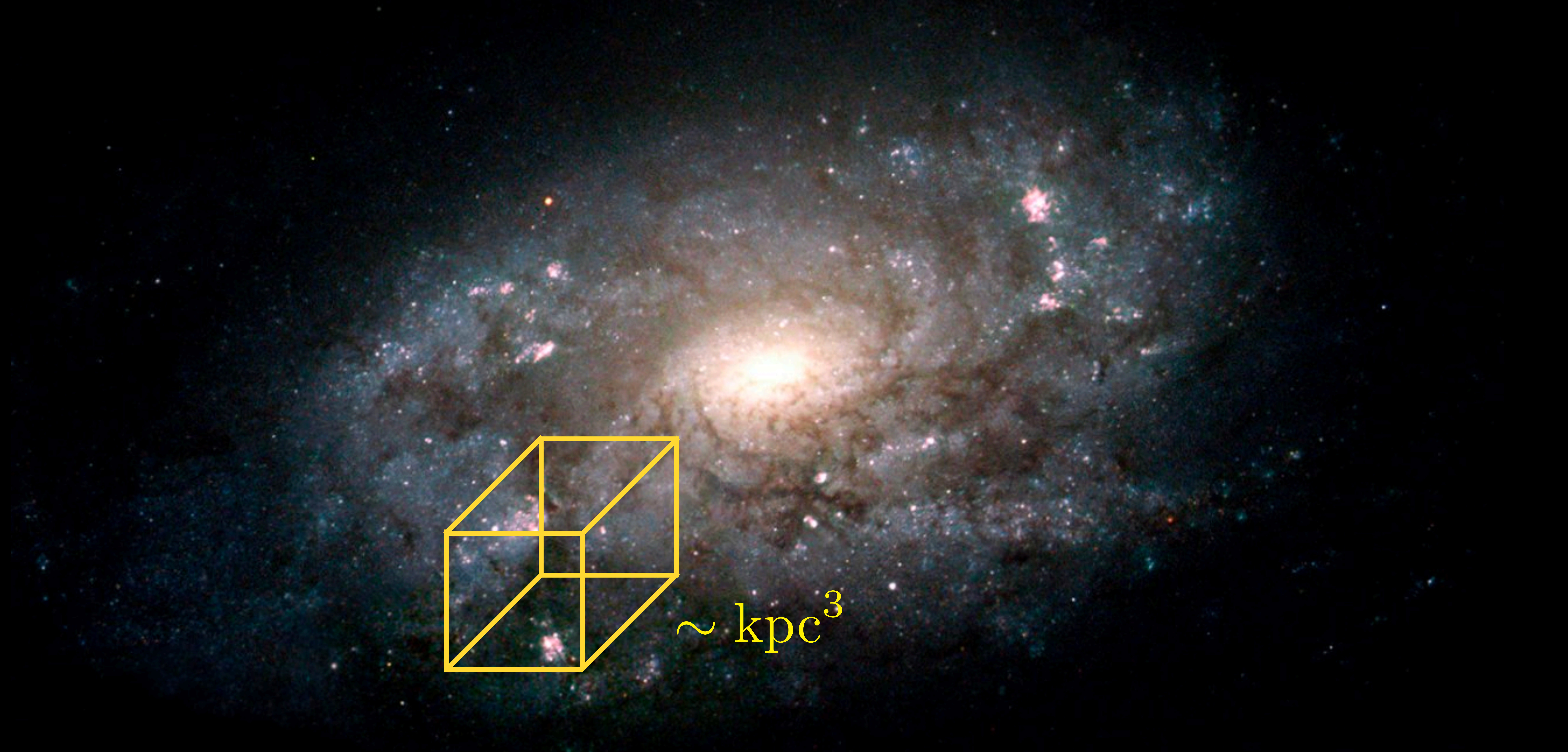
See also [Dror and Verner 2406.03526]

Remark I

all of the results shown here are sensitive to

ULDM density around/within the solar system

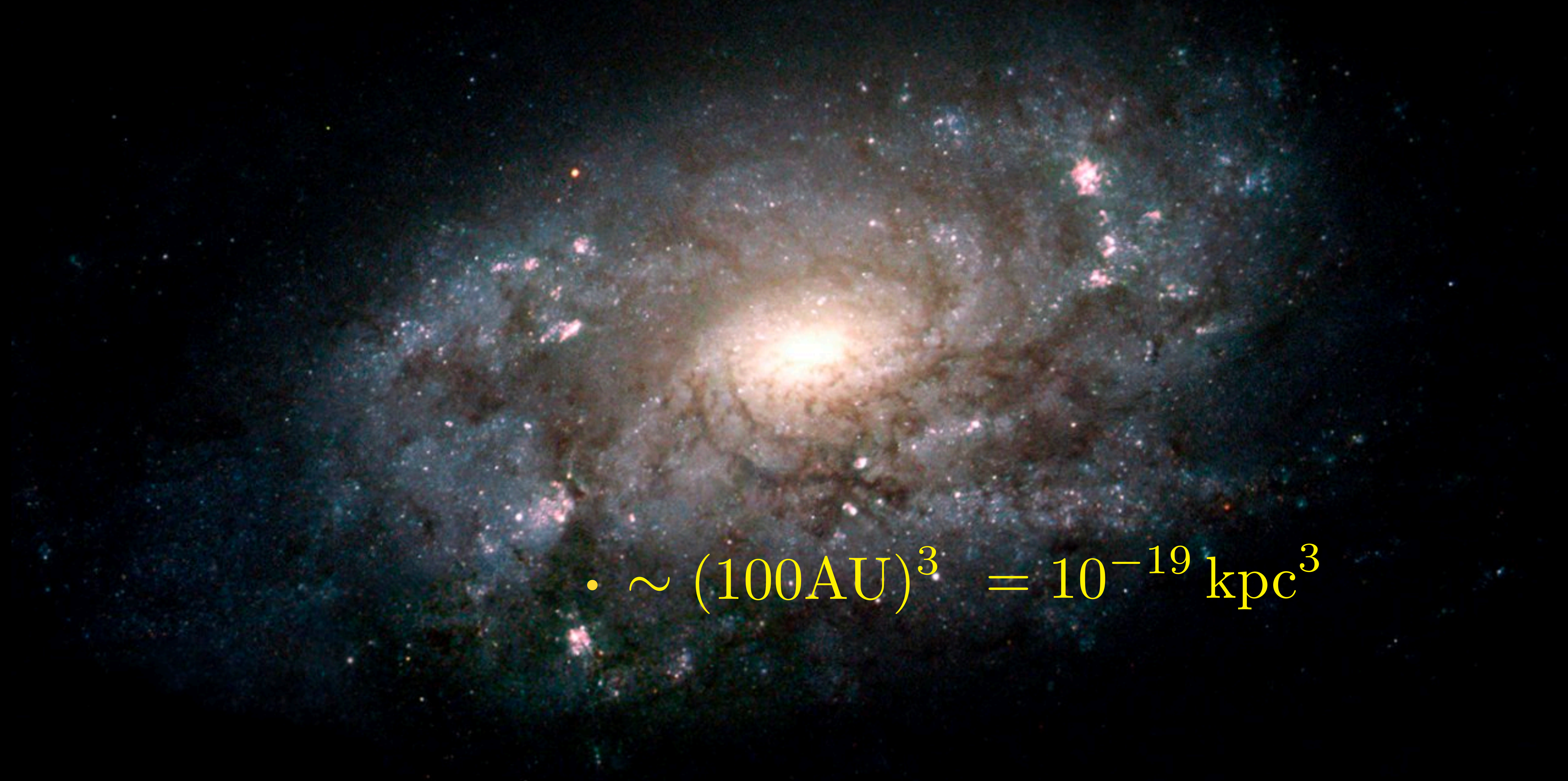
local dark matter density is often derived over kpc scales



$$\rho_0 = 0.4\,\mathrm{GeV/cm^3}$$

is an ***average density over the volume of kpc*** what we are probing is

(or what matters for all terrestrial DM detector is)



- $\sim (100\mathrm{AU})^3 \;\; = 10^{-19}\,\mathrm{kpc}^3$

currently no measurement on this scale exists only constraints exist

$$\rho/\rho_0 \lesssim 10^{11}$$

From geodetic satellite and LLR
[Adler (08)]

$$\rho/\rho_0 \lesssim 6 \times 10^{6}$$

From asteroids in the solar system
[Tsai, Eby et al (22)]

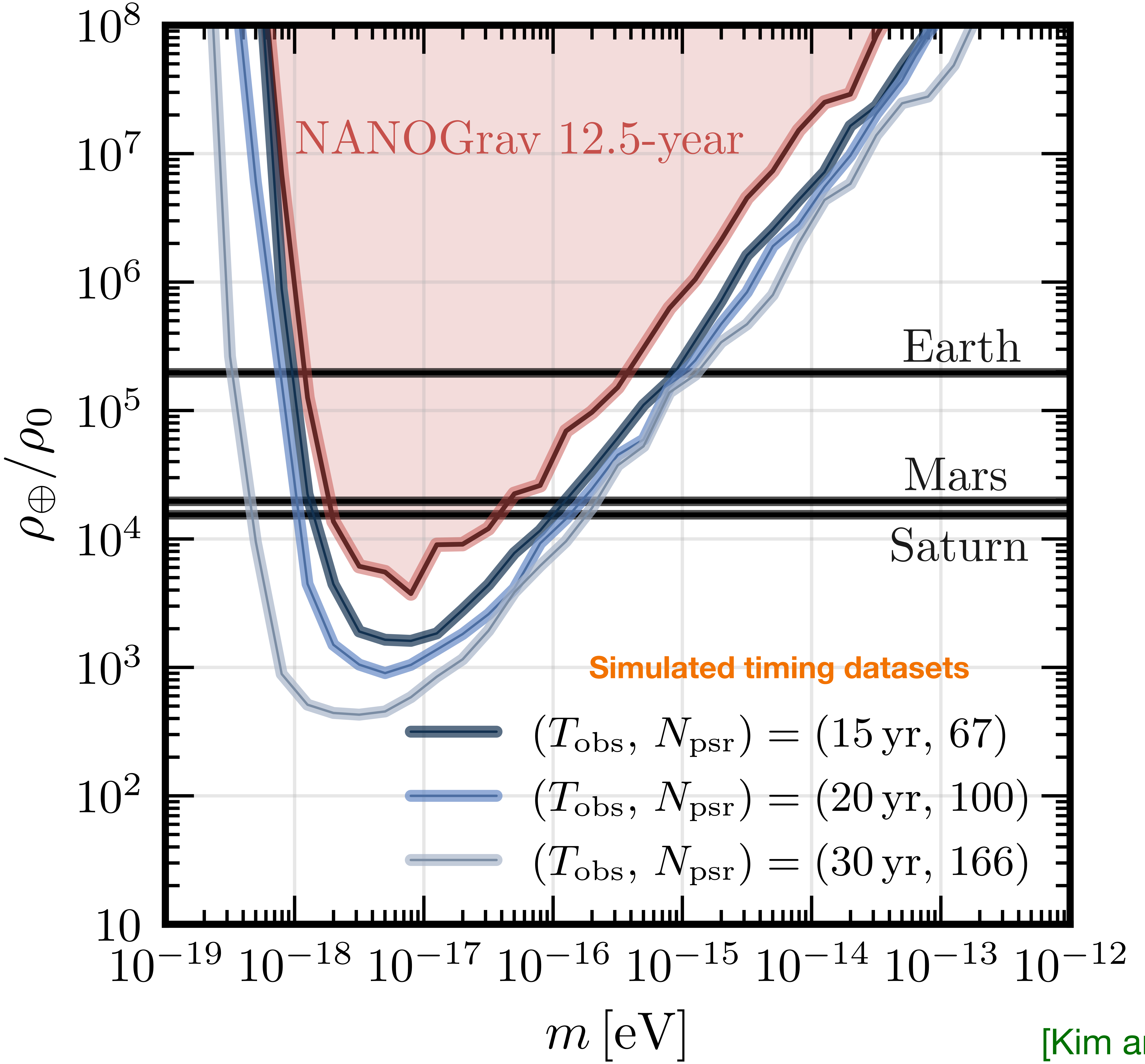
$$\rho/\rho_0 \lesssim 2 \times 10^{4}$$

From solar system ephemerides
[Pitjev, Pitjeva (13)]

GW detectors will provide
one of the strongest probes of ULDM density
within/around the solar system

Remark II

prospects of pulsar timing array



[Kim and Mitridate 2312.12225]

quite realistic choice of simulated data

$$\sigma_{\text{TOA}} \sim 500\,\text{ns}$$

$$\log_{10} A_a \sim \mathcal{N}(-16, 1)$$

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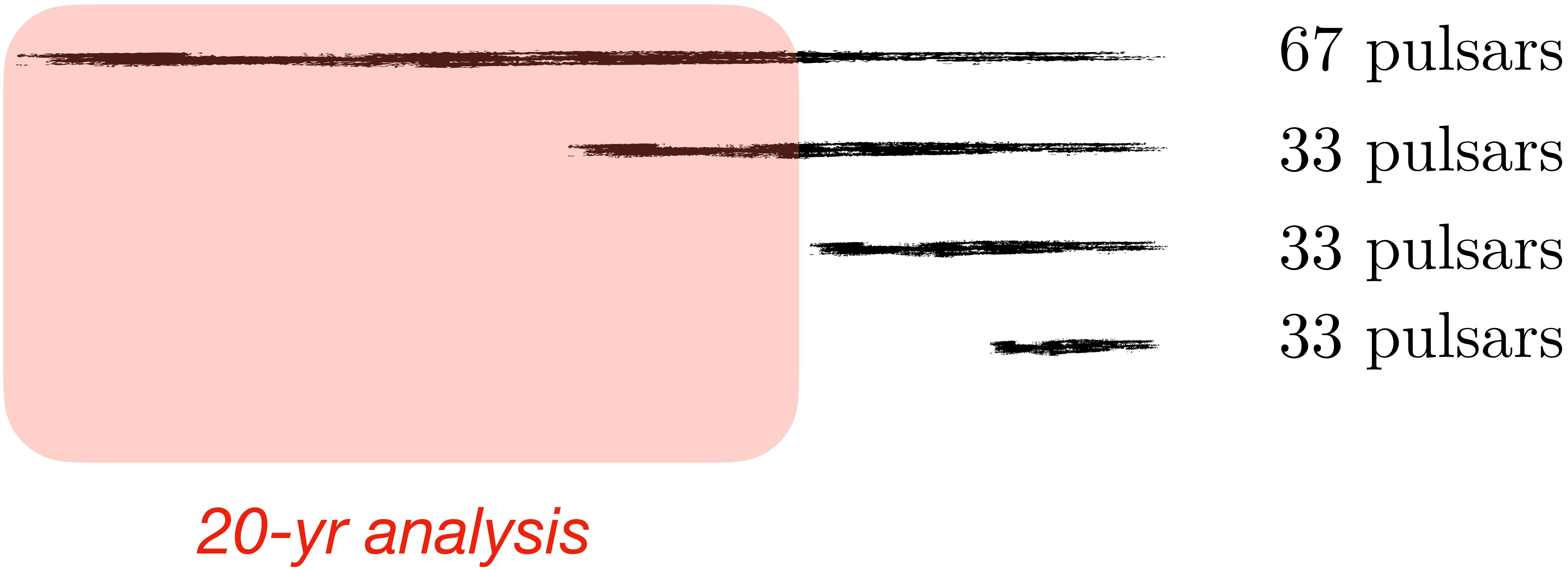

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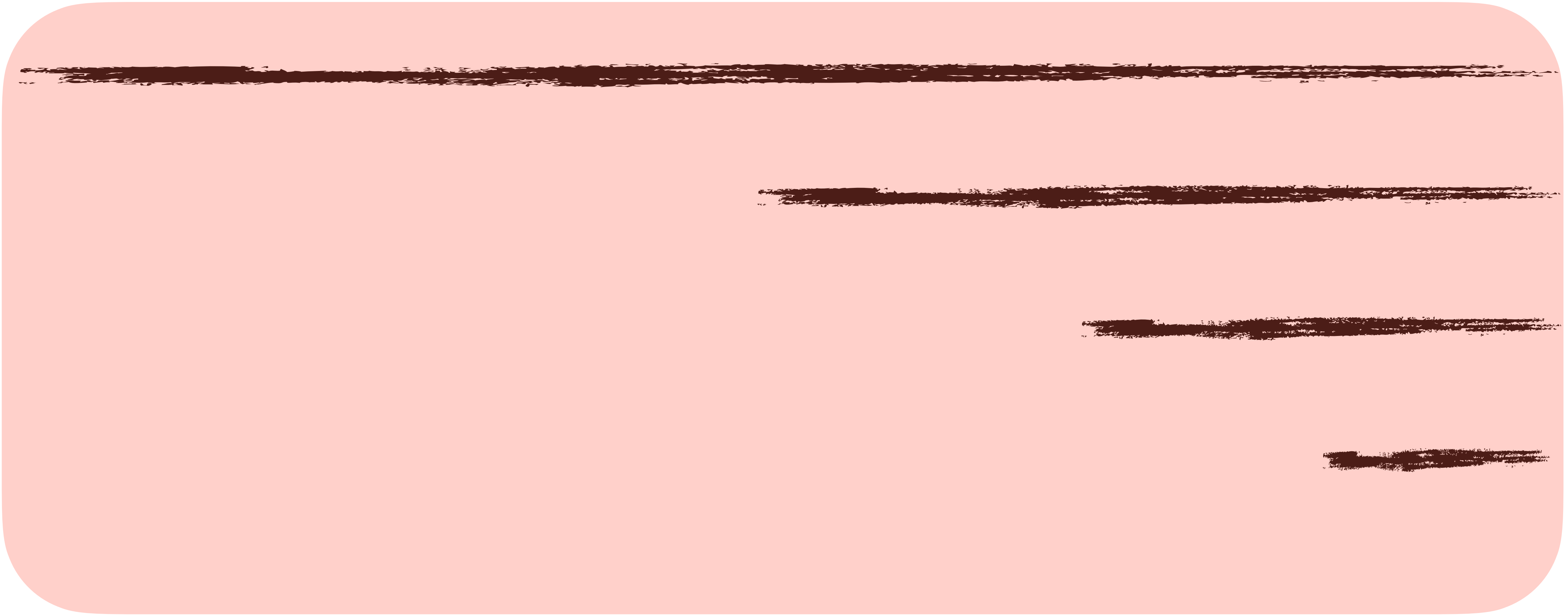


67 pulsars

33 pulsars

33 pulsars

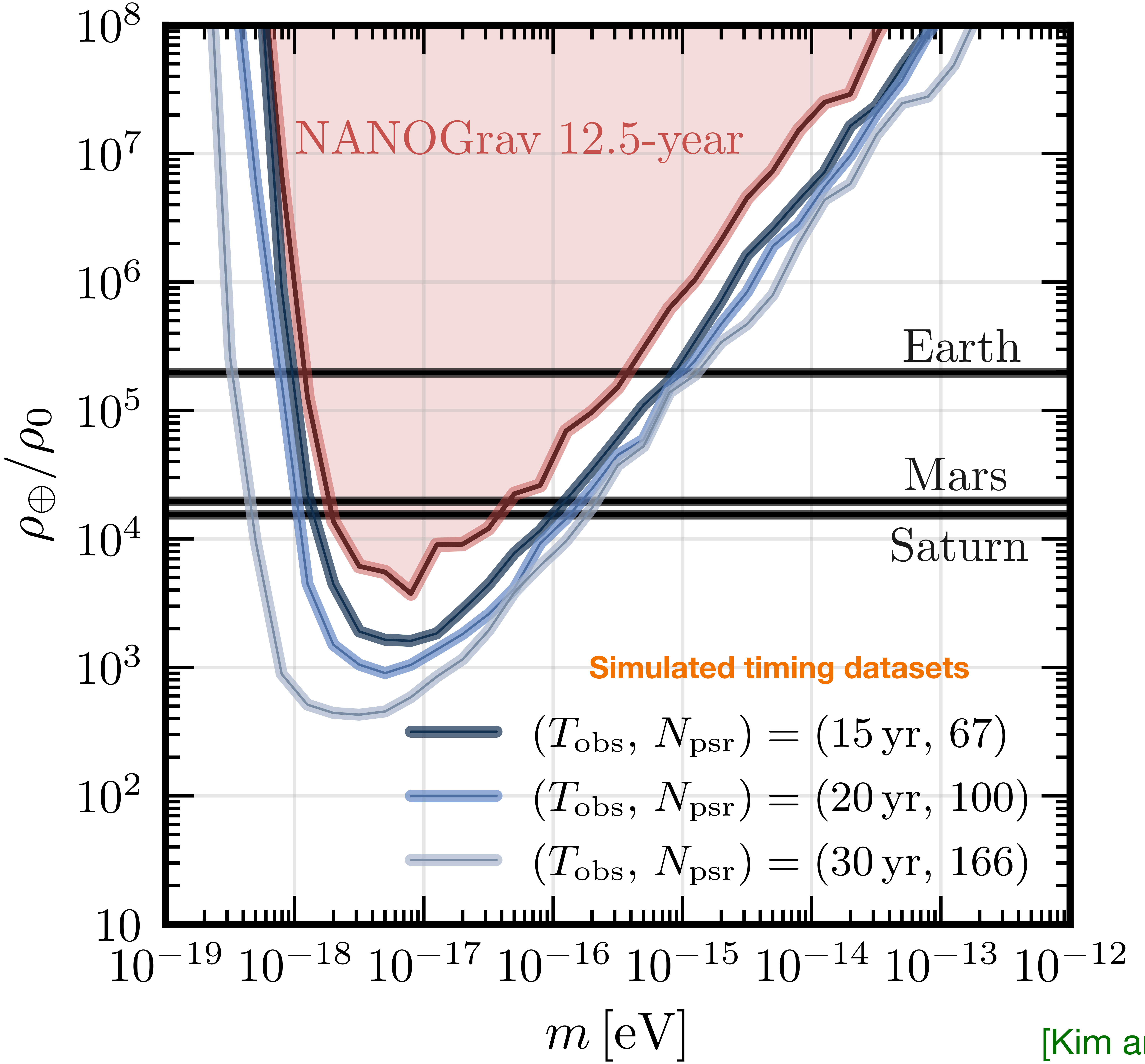
33 pulsars

30-yr analysis

Next International Pulsar Timing Array (IPTA) data release
is expected to include

$$N_{\rm psr} \sim 100$$

$$T_{\rm obs} = 20^{+}\,{\rm yr}$$



[Kim and Mitridate 2312.12225]

with next-gen ratio telescope (e.g. Square Kilometer Array)
an order of magnitude or more improvement might be feasible

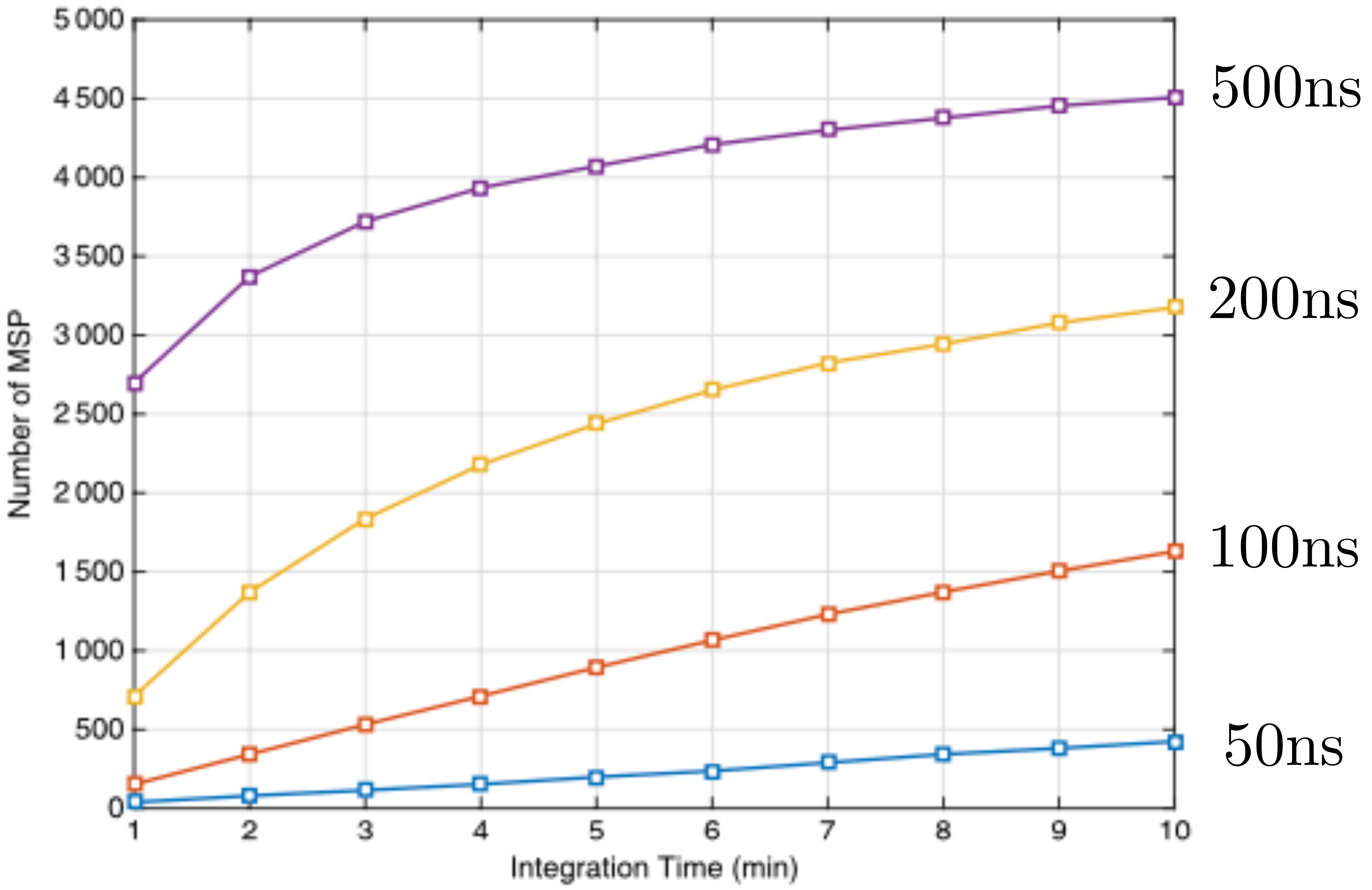


Figure 10. Numbers of MSPs that can archive a certain RMS noise level (or better) with varying integration time. Colour lines indicate different RMS noise levels (from bottom to top): 50 ns (blue), 100 ns (red), 200 ns (yellow), and 500 ns (purple).

[SKA, Weltman et al (20)]