

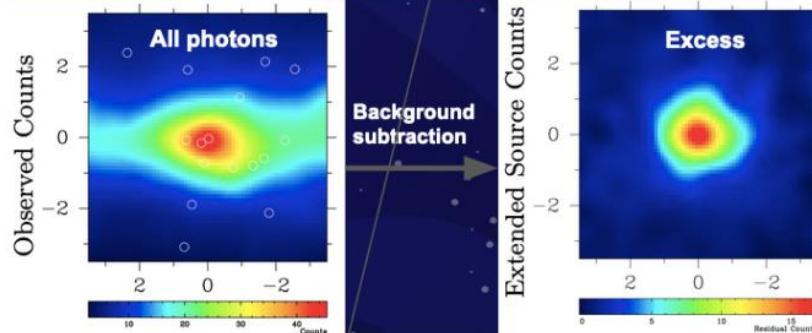
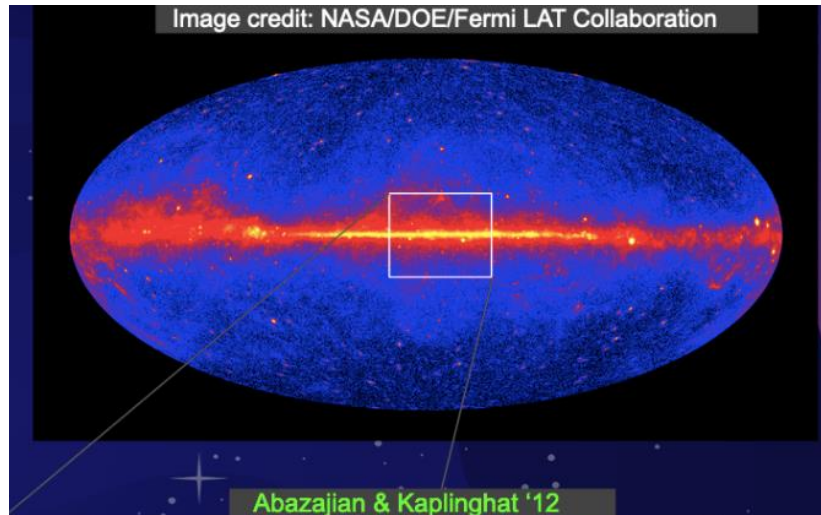
Probing the pulsar explanation of the GeV excess using GW

Yue Zhao

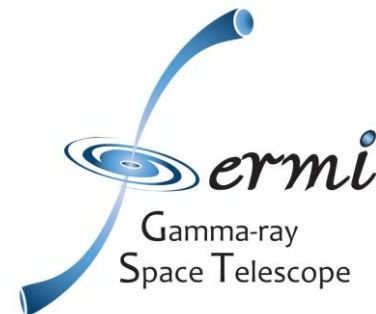
University of Utah, Salt Lake City



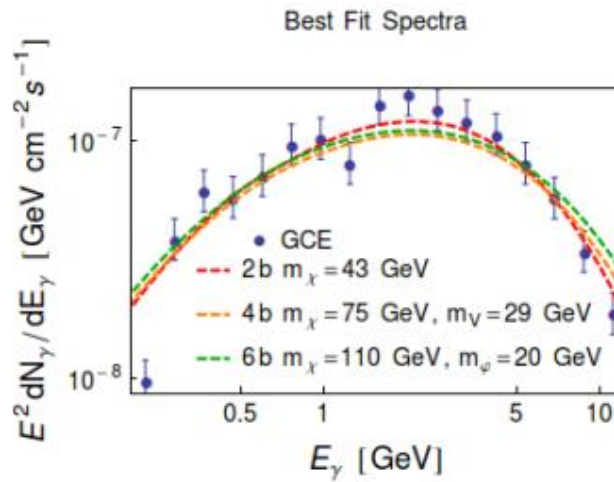
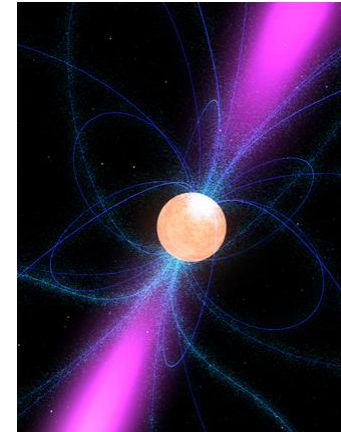
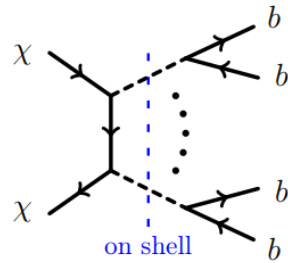
A GeV excess at the Galactic Center:



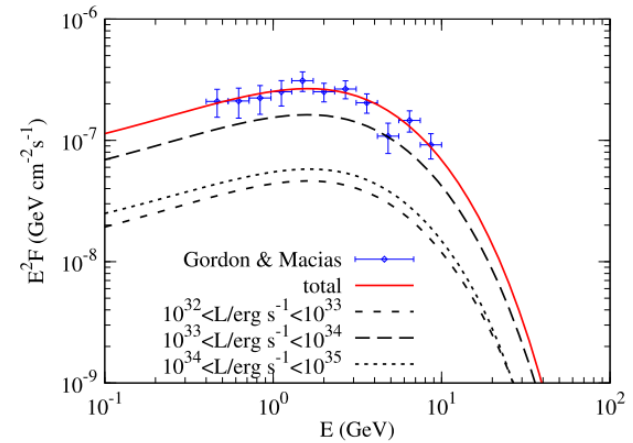
Well motivated.
Correct relic abundance.
Searched for decades.



Two explanations:



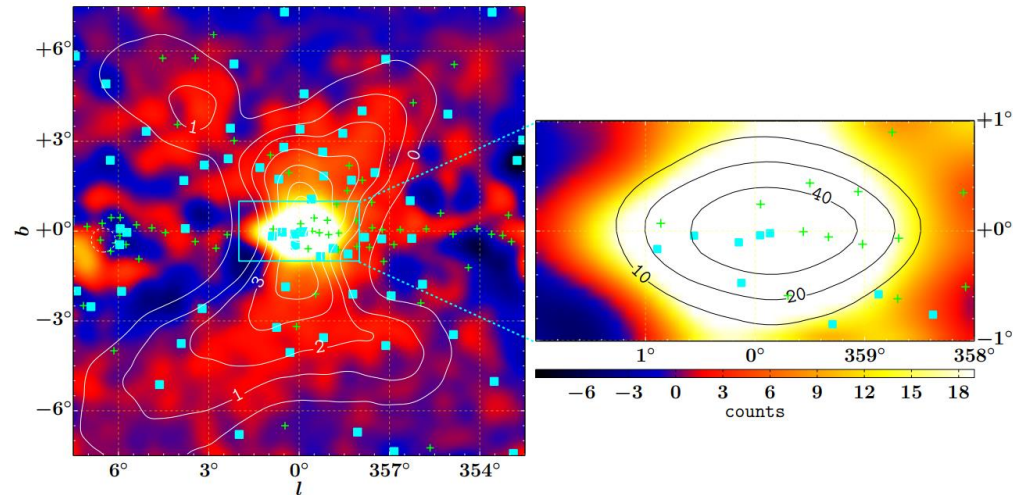
Abdullah, et. al.
 Phys. Rev. D 90, 035004 (2014)



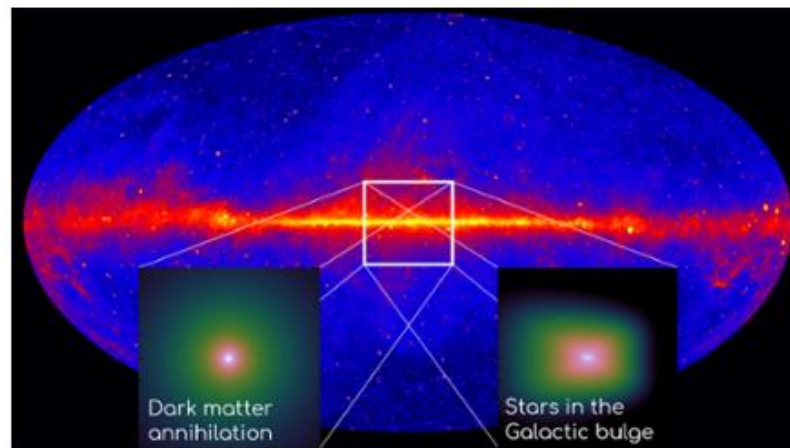
Yuan, et. al.
 JHEAp 3 (2014) 1

Efforts to distinguish these two explanations:

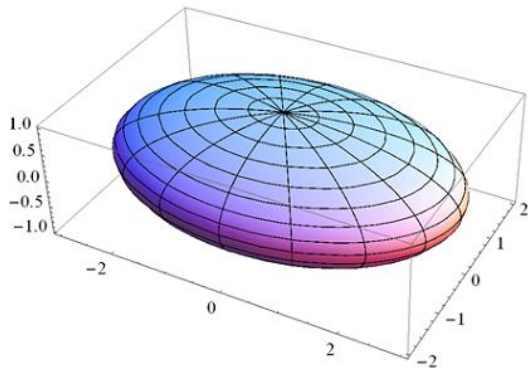
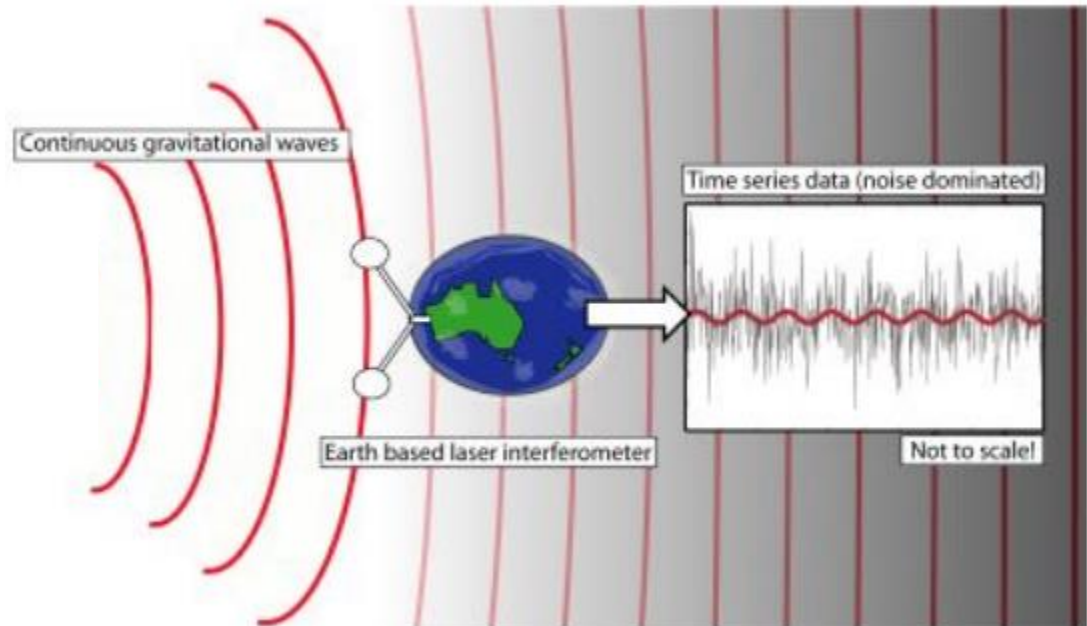
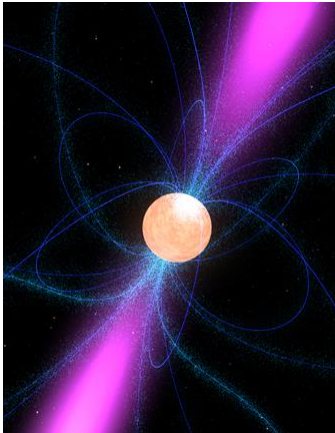
- Smoothness: Point Source v.s. Smearred Distribution



- Morphology: Spherical v.s. Bulge-like



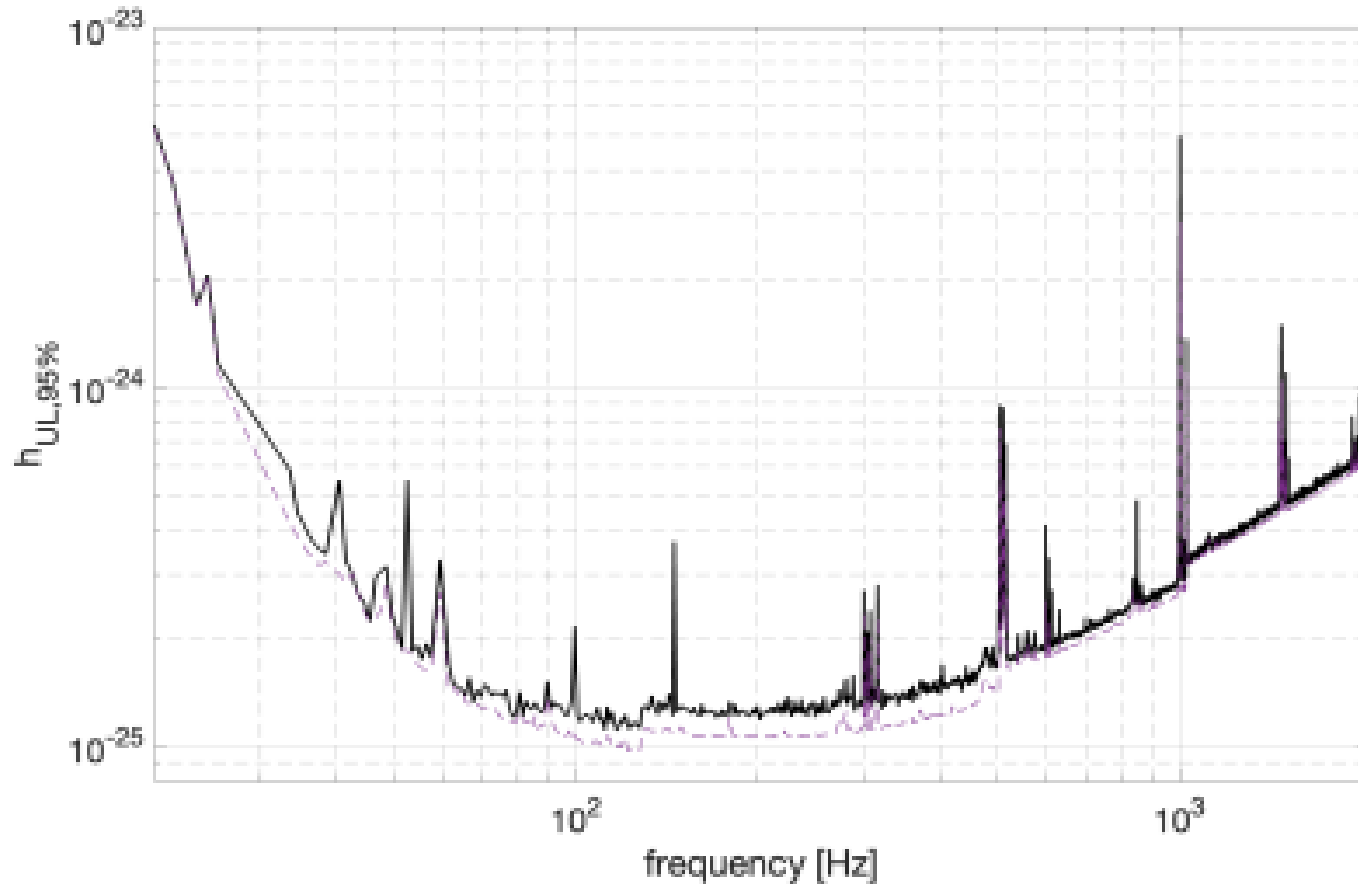
GW channel can be useful:



$$h_0 = \frac{16\pi^2 G}{c^4} \frac{I_{zz} \epsilon f_{\text{rot}}^2}{d}$$

WIMP DM:

The LVK collaboration
Phys. Rev. D 106, 102008

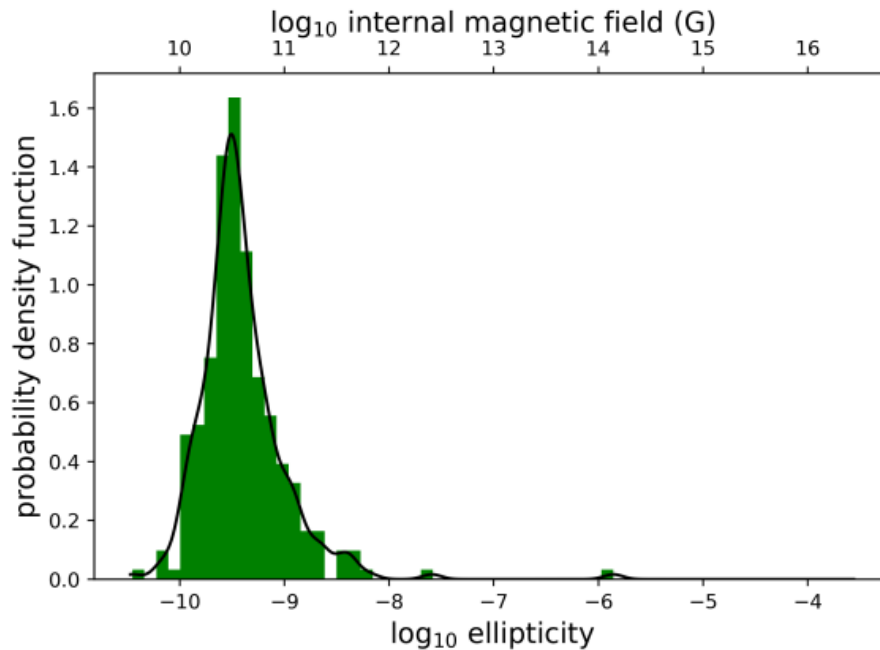


Ellipticity distribution:

$$h_0 = \frac{16\pi^2 G}{c^4} \frac{I_{zz} \epsilon f_{\text{rot}}^2}{d}$$

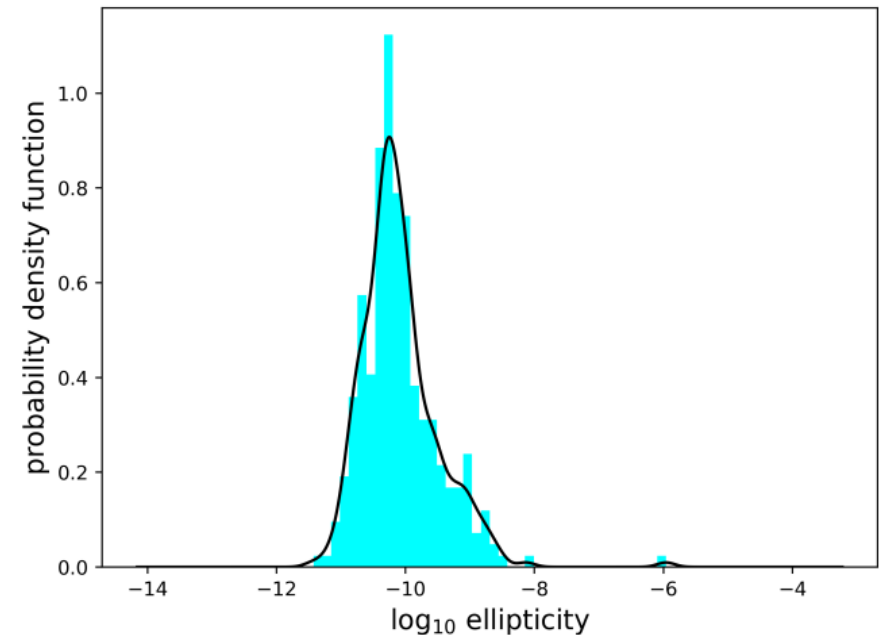
$$\epsilon \equiv \frac{|I_{xx} - I_{yy}|}{I_{zz}}$$

principal moments of inertia



$$\epsilon \approx 10^{-8} \left(\frac{B_{\text{int}}}{10^{12} \text{Gs}} \right)$$

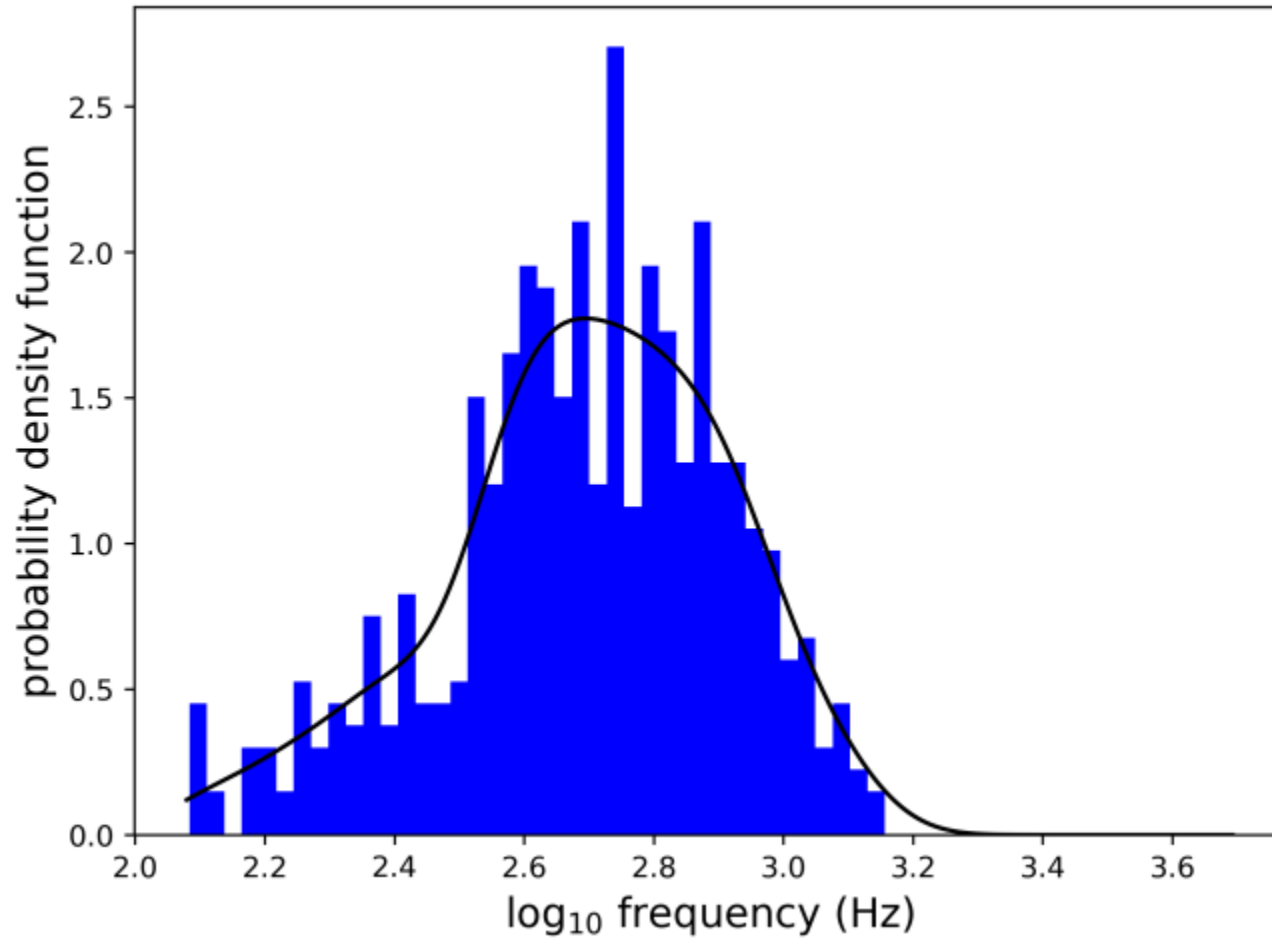
$$B_{\text{int}} = 150 B_{\text{ext}}$$



GW radiation accounts for 1% rotational energy loss.

Frequency distribution:

$$h_0 = \frac{16\pi^2 G}{c^4} \frac{I_{zz} \epsilon f_{\text{rot}}^2}{d}$$

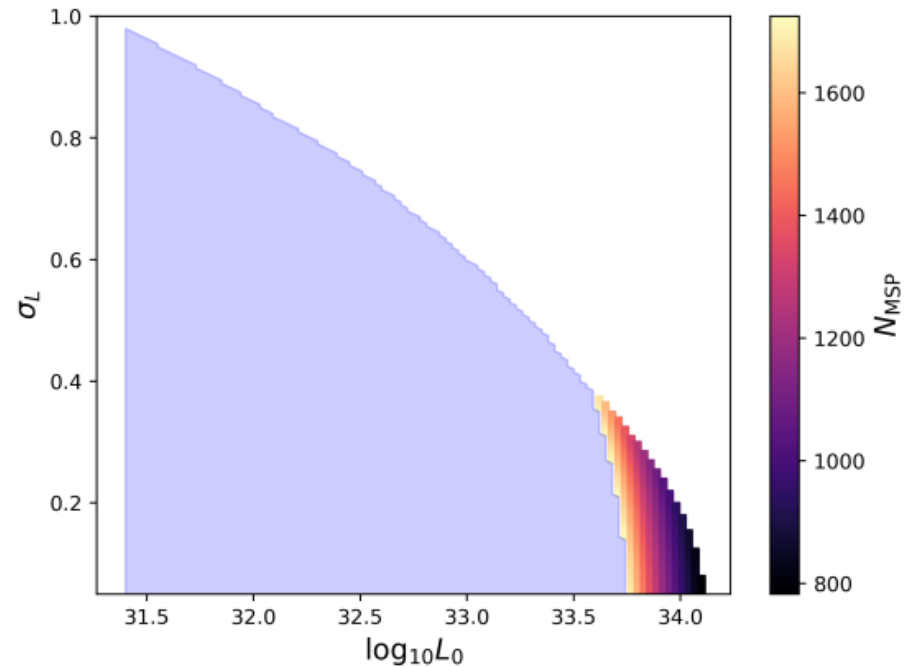
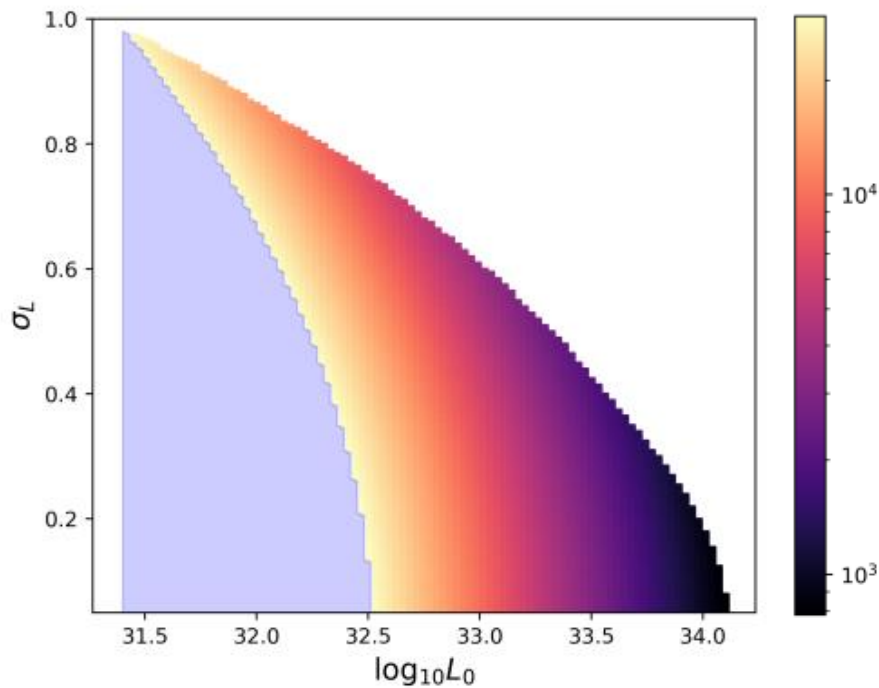


WIMP DM:

Andrew Miller, Y.Z.

Phys.Rev.Lett. 131 (2023) 8, 081401

$$\frac{dP(L)}{dL} = \frac{\log_{10} e}{\sigma_L \sqrt{2\pi} L} \exp\left(-\frac{\log_{10}^2(L/L_0)}{2\sigma_L^2}\right)$$



1% energy loss through GW

$d = 8$ kpc

$I_{zz} = 10^{38} \text{kg}\cdot\text{m}^2$

$B_{\text{int}} = 150 B_{\text{ext}}; d = 8$ kpc.

$I_{zz} = 5 \times 10^{38} \text{kg}\cdot\text{m}^2$

Conclusion:

The origin of the GeV excess at the galactic center:

dark matter or milli-second pulsars

We need a lot of these MSPs to explain the GeV excess.

GW provides a good probe to distinguish these two interpretations.

The current limits from LVK is already quite strong.

We expect more parameter space can be probed in near future.

Better sensitivity in the next LVK run.

Optimized analysis pipeline is under construction.

SGWB vs Continuous wave.