Exploring DSNB boosted sub-GeV dark matter: insights from XENONnT and LZ experiments

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Table of contents

- 1 Overview
- 2 Introduction
- 3 Dark matter landscape
- 4 Thermal relic dark matter vs boosted dark matter
- 5 DSNB boosted dark matter
 - DSNB
 - Boosted dark matter flux at the underground detectors
 - Dark matter signal at the underground detectors
 - Resulting Limits

6 Conclusions



Introduction



Motivation for searching the Dark Matter (DM)

- Cold DM: It is a non-luminous matter which occupies 27% of the mass and energy in the observable Universe. It does not interact with photons and interacts only "weakly" with ordinary matter.
- Astronomical and cosmological observations at various scales:
 - (i) Rotation curves of spiral galaxies and galaxy clusters
 - (ii) Gravitational lensing
 - (iii) Cosmic Microwave background (CMB) fluctuations



- Direct Detection Experiments: XENONnT, LUX-ZEPLIN, Super-CDMS, Dark-Side, PandaX-4T, etc.
- Indirect Detection Experiment: IceCube, HESS, MAGIC, etc.
- Accelerator searches: ATLAS, CMS at CERN

DM landscape: a wide mass range





DM landscape: a wide mass range









Thermal relic DM vs boosted DM



The maximum recoil energy of the target:

 $T_r^{\max} pprox rac{Q^2}{2m_T} pprox rac{2m_\chi^2 m_T v_\chi^2}{(m_\chi + m_T)^2}$

Thermal relic DM vs boosted DM



The maximum recoil energy of the target:



Thermal relic DM vs boosted DM



The maximum recoil energy of the target:



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DSNB Boosted Dark Matter

15th October 2024 4 / 13

DSNB boosted dark matter

Diffuse Supernova Neutrino Background (2) IISER Bhopal



Right after the first star formation event, the Universe has been surrounded by an isotropic flux of MeV-energy neutrinos and antineutrinos of all flavors, produced from all supernovae events from the core-collapse explosions of huge stars throughout the Universe. This cumulative and isotropic flux of MeV neutrinos form DSNB.



BDM Flux At The Underground Detectors

The DSNB-boosted DM differential flux,

$$rac{d\Phi_{\chi}}{dT_{\chi}} = D_{ ext{halo}} \sum_{lpha} \int_{E_{
u}^{ ext{min}}}^{E_{
u}^{ ext{max}}} dE_{
u} rac{1}{m_{\chi}} rac{d\sigma_{
u_{\chi}}}{dT_{\chi}} rac{d\Phi_{
u_{lpha}}^{ ext{DSNB}}}{dE_{
u}}$$

D_{halo} encodes the line of sight integral of DM density within our galactic halo,

$$D_{
m halo} = \int_{\Delta\Omega} rac{d\Omega}{4\pi} \int_0^{\ell_{
m max}}
ho_{
m NFW}[r(\ell,\psi)] d\ell$$

We consider Navarro-Frenk-White (NFW) profile for galactic DM density,

$$\rho_{\rm NFW}(r) = \rho_{\odot} \left[\frac{r}{r_{\odot}}\right]^{-1} \left[\frac{1+\frac{r_{\odot}}{r_{\rm s}}}{1+\frac{r}{r_{\rm s}}}\right]^2$$

For the differential cross section, we assume the constant cross section approximation, i.e.

$$rac{d\sigma_{
u\chi}}{dT_{\chi}} = rac{\sigma_{
u\chi}}{T_{\chi}^{
m max}}$$

BDM Flux At The Underground Detectors

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u} rac{1}{m_{\chi}} rac{d\sigma_{
u_{\chi}}}{dT_{\chi}} rac{d\Phi_{
u_{lpha}}^{
m DSNB}}{dE_{
u}}$$

DM flux gets attenuated by the elements of the atmosphere and Earth before reaching to the underground detector. However as $n_i^{\text{atm}} << n_i^{\text{Earth}}$, we have **neglected the attenuation due to elements of the atmosphere**.



$$\frac{d\sigma_{\chi i}}{dT_i} = \frac{\sigma_{\chi i}}{T_i^{\max}}$$
$$i \equiv \{e, \mathcal{N}\}$$

BDM Flux At The Underground Detectors (2) IISER Bhopal

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DM flux gets attenuated by the elements of the atmosphere and Earth before reaching to the underground detector. However as $n_i^{\text{atm}} << n_i^{\text{Earth}}$, we have **neglected the attenuation due to elements of the atmosphere**.



6/13

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Implications of nuclear form factor





DM signal at the underground detectors



After reaching the underground detector, the DSNB-boosted DM can scatter off both the electrons and nuclei of the target material, triggering both electronic and nuclear recoils. The differential event rate with respect to the recoil energy T_i can be written as,





Effect of Earth attenuation in the resulting limits





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DSNB Boosted Dark Matter

15th October 2024 9 / 13

Resulting Limits





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Conclusions



- SNB Boosted DM produces a subdominant, semi-relativistic component of Galactic DM.
- © Consideration of Earth attenuation is crucial for accurate interpretation of experimental results.
- Although a significant part of our constraints lie in a region of parameter space already probed by other searches, these results highlight the complementarity and significance of the LZ and XENONnT data in probing the sub-GeV DM parameter space.

THANK YOU



χ^2 function utilized



For the analysis of LZ data, we have performed a spectral analysis using the following Poissonian χ² function

$$\begin{split} \chi^{2}(\overrightarrow{\mathcal{S}};\alpha,\beta,\delta) = & 2\sum_{i=1}^{51} \left[R_{\text{pred}}^{i}(\overrightarrow{\mathcal{S}};\alpha,\beta,\delta) - R_{\text{exp}}^{i} + R_{\text{exp}}^{i} \ln\left(\frac{R_{\text{exp}}^{i}}{R_{\text{pred}}^{j}(\overrightarrow{\mathcal{S}};\alpha,\beta,\delta)}\right) \right] + \\ & \left(\frac{\alpha}{\sigma_{\alpha}}\right)^{2} + \left(\frac{\beta}{\sigma_{\beta}}\right)^{2} + \left(\frac{\delta}{\sigma_{\delta}}\right)^{2} \,, \end{split}$$

• The following Gaussian χ^2 function is used for the analysis of XENONnT data

$$\chi^{2}(\overrightarrow{\mathcal{S}};\boldsymbol{\beta}) = \sum_{i=1}^{30} \left(\frac{R_{\mathsf{pred}}^{i}(\overrightarrow{\mathcal{S}};\boldsymbol{\beta}) - R_{\mathsf{exp}}^{i}}{\sigma^{i}} \right)^{2} + \left(\frac{\boldsymbol{\beta}}{\sigma_{\boldsymbol{\beta}}} \right)^{2}$$

Geophysical properties of Earth



We model the Earth's interior as a sphere of constant electron and nuclear densities ($n_e = 8 \times 10^{23} \text{ cm}^{-3}$ and $n_N = 3.44 \times 10^{22} \text{ cm}^{-3}$), based on the abundances of the main elements as shown in following table.

Element	Mass Number (A)	Relative Abundance (%)	$n_\mathcal{N}~(cm^{-3})$
Fe	55.845	32.1	$6.11 imes 10^{22}$
0	15.999	30.1	$3.45 imes10^{22}$
Si	28.086	15.1	$1.77 imes10^{22}$
Mg	24.305	13.9	$1.17 imes 10^{22}$
S	32.065	2.9	$2.33 imes10^{21}$
Ca	40.078	1.5	$7.94 imes10^{20}$
AI	26.982	1.4	1.09×10^{21}