Wave optics lensing of GW :

Theory and phenomenology of triple systems in the LISA band

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In collaboration with Giulia Cusin, Cyril Pitrou, Jean-Philippe Uzan

arXiv:2404.07186 , published in Phys. Rev. D

UNIVERSITÉ DE GENÈVE

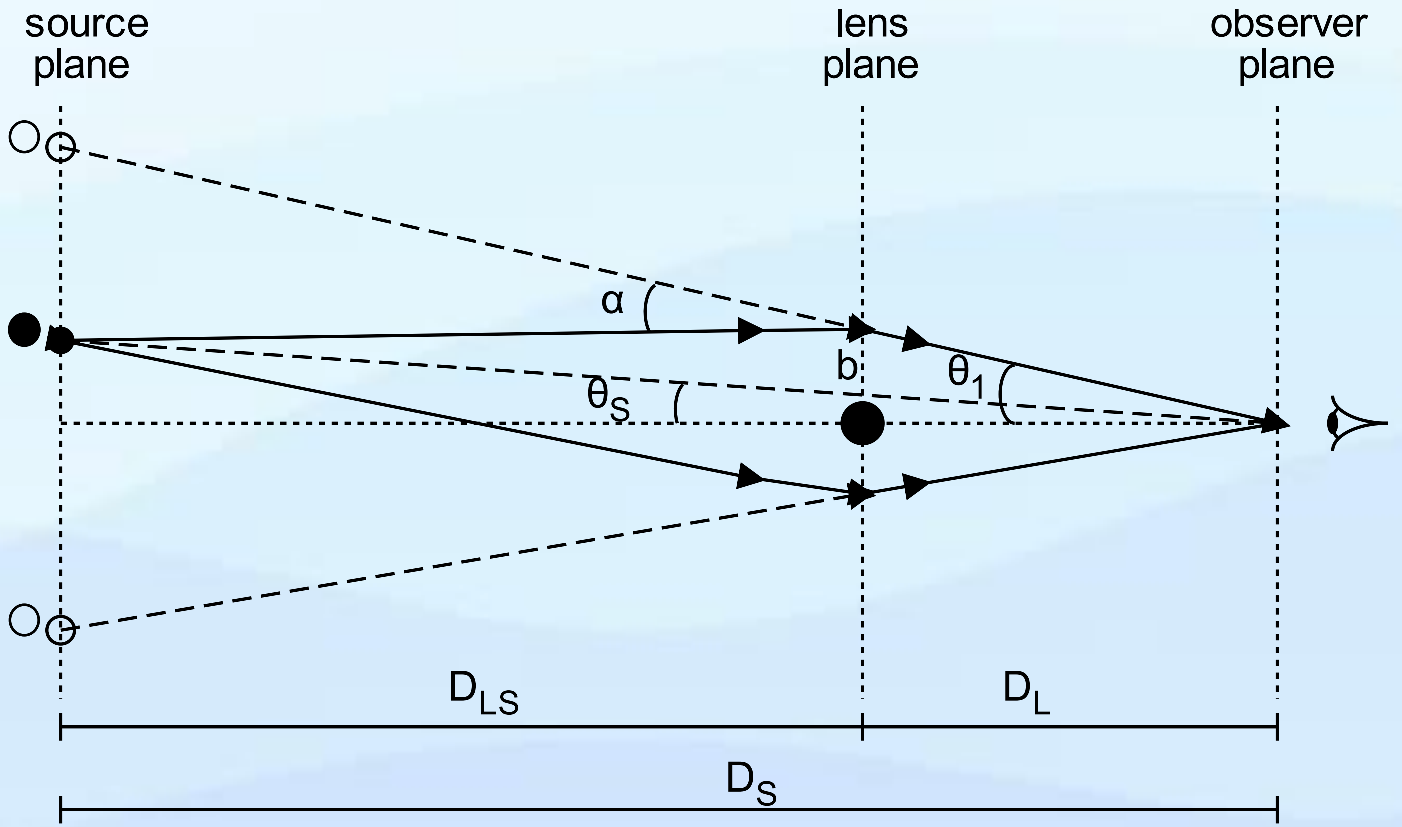
Outline

1. Gravitational lensing : geometric vs wave optics
2. GW lensing, wave optics in the *usual* way
3. **Our contribution** : GW lensing in wave optics considering the tensorial structure
4. Application to LISA-band hierarchical triple systems

GW lensing: geometrical optics

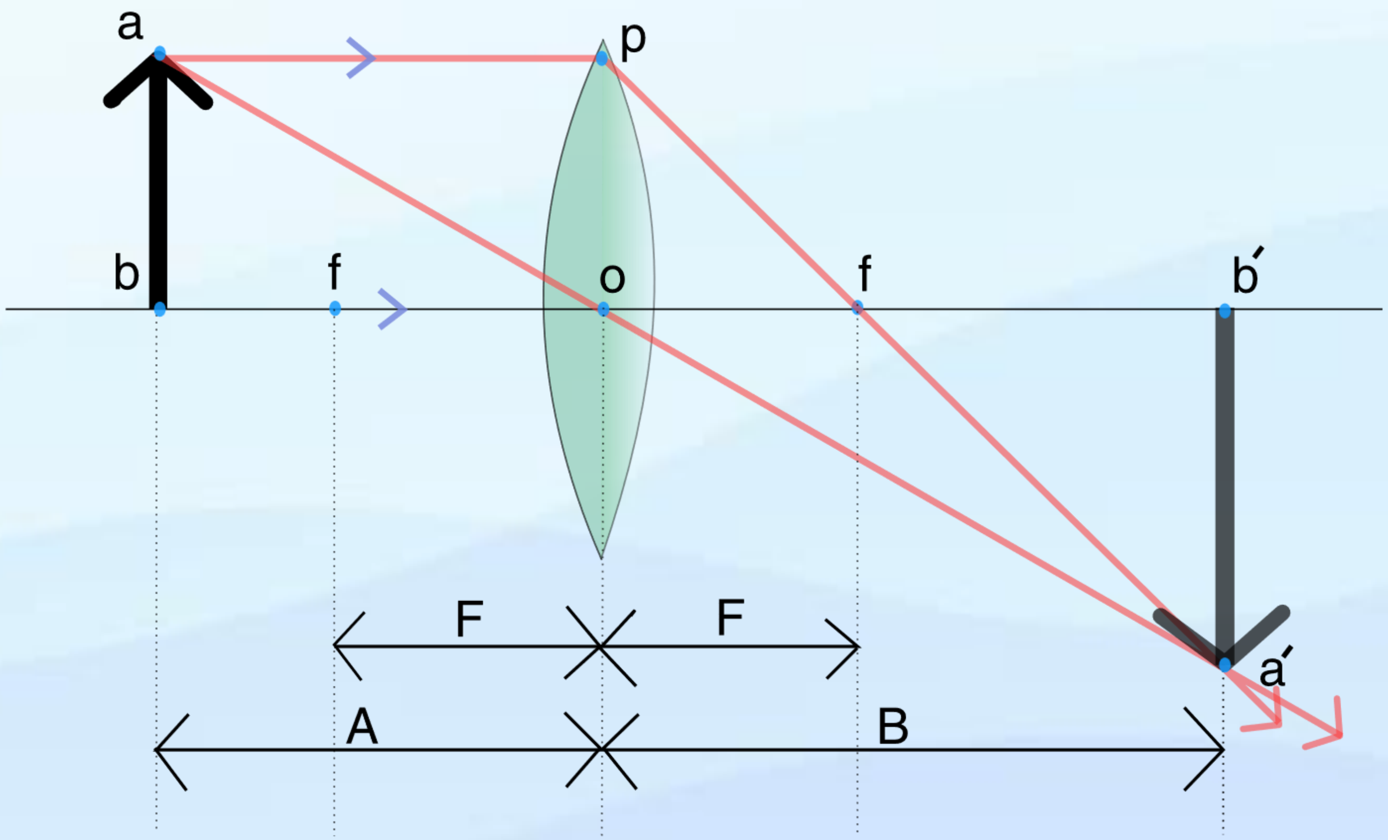
Assume you have a notion of ***ray*** :

→ Usual lensing picture (deflection angle, etc.)



Gravitational lensing: geometrical optics

Conceptually similar to (my) undergraduate lab optics



GW lensing: wave optics

But just as EM radiation, GW are ... **waves** !

Geometrical optics is just a high frequency approximation, which **breaks down** when

$$\boxed{\lambda_{\mathrm{wave}} \gtrsim \text{(lens size)}\,.}$$

At the fundamental level, signal obeys a wave equation

$\rightarrow$ allows for diffraction, interference, ...

Wave optics features

Criterion $\boxed{\lambda_{\mathrm{wave}} \gtrsim \text{(lens size)}}$

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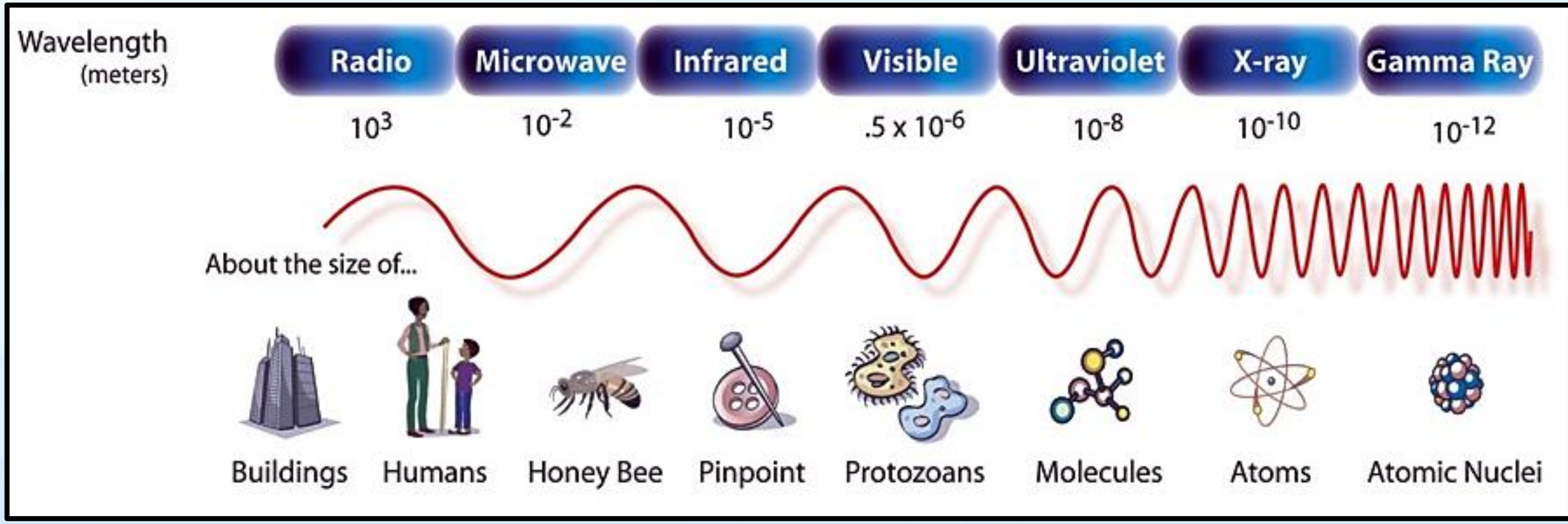
EM spectrum:



Credits : NASA

Wave optics features

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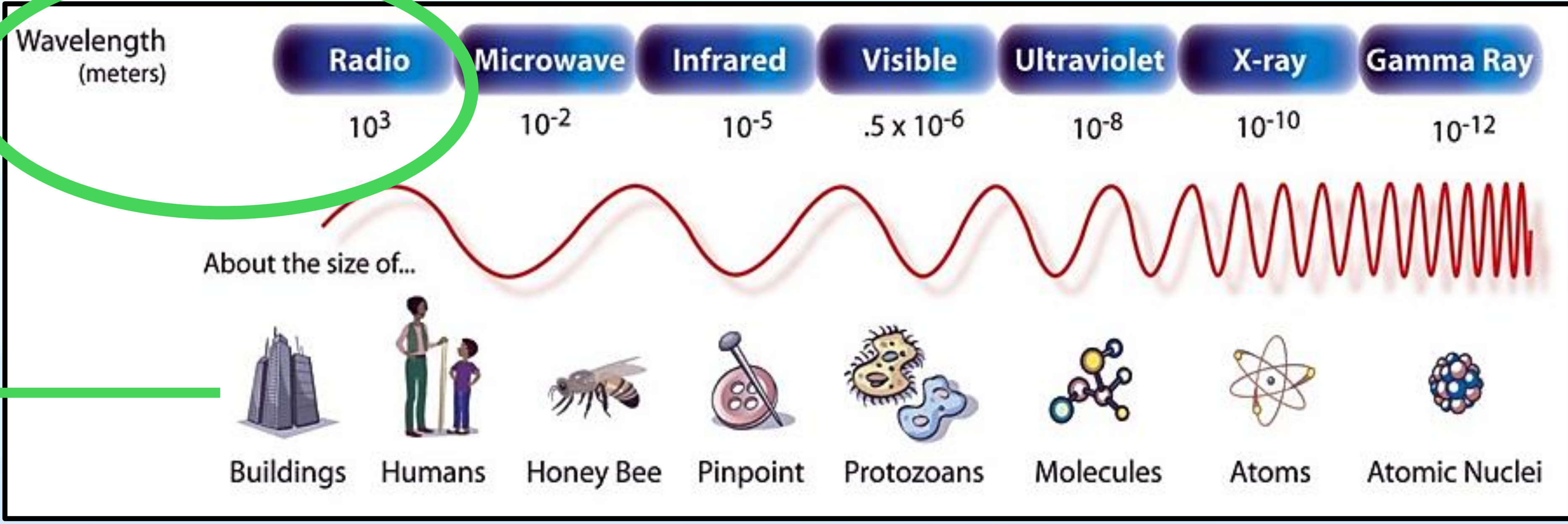

Credits : NASA

Rather irrelevant for astro/cosmo

→ **Wave effects very subdominant in practice**

GW lensing: wave optics

LISA has a best sensitivity around 10^{-3}Hz.

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At this frequency, **low mass AGN** with $M \sim \mathcal{O}(10^6) M_\odot$ fulfil the wave optics requirement

$$\lambda_{GW} > \frac{2GM}{c^2}$$

Equivalently : $\omega M < 1$ (natural units)

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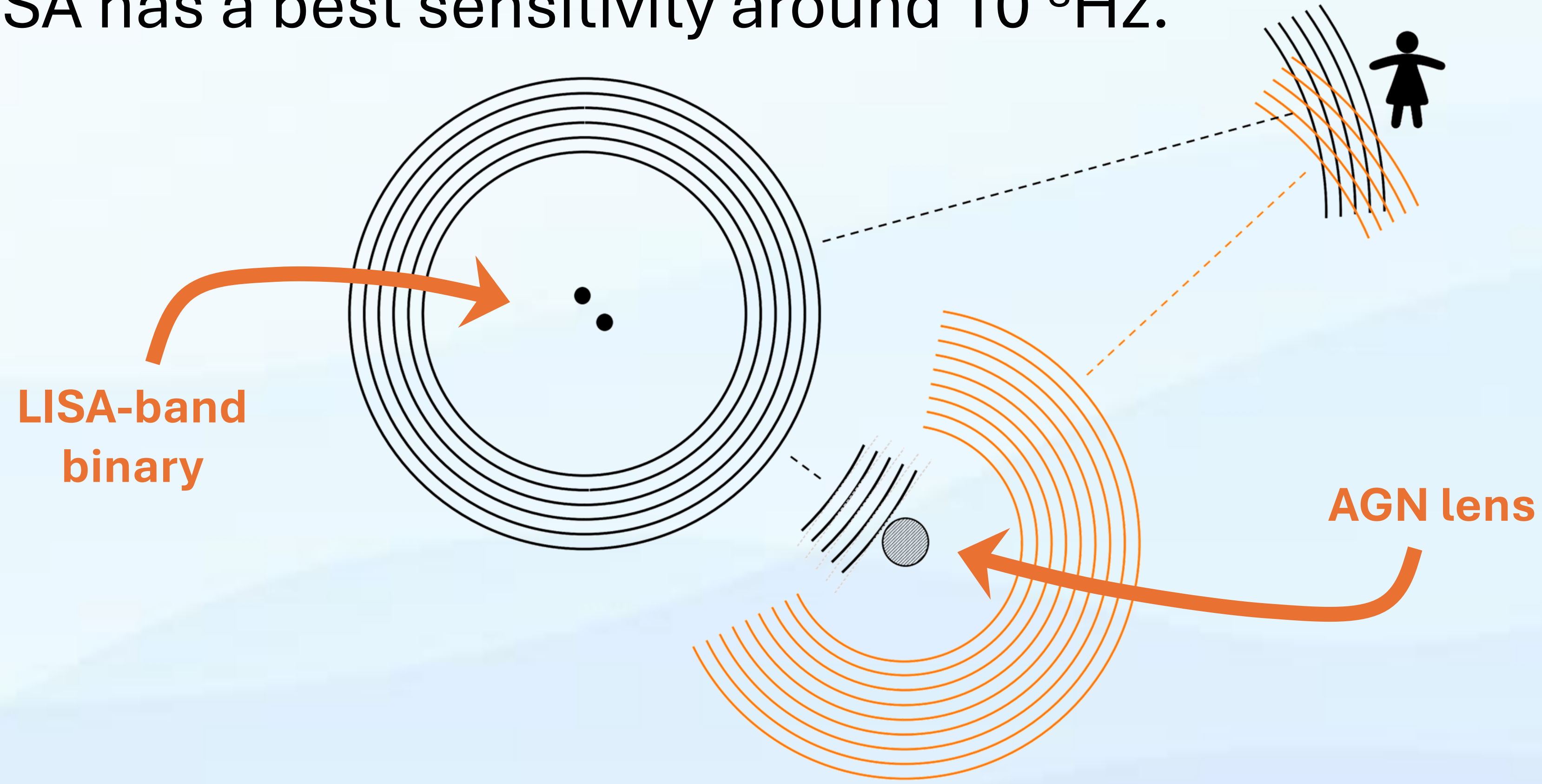

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Equivalently : $\omega M < 1$ (natural units)

How to capture wave effects ?

GW lensing: wave optics

Start with : $\mathrm{d}s^2 = (\bar{g}_{\mu\nu} + h_{\mu\nu})\mathrm{d}x^\mu \mathrm{d}x^\nu$

E.g. gauge fixing : $h^{\nu}_{\mu;\nu} = 0, \quad h^{\mu}{}_{\mu} = 0$

→ Wave equation :

$$h_{\mu\nu;\alpha}{}^{;\alpha} + 2\bar{R}_{\alpha\mu\beta\nu}\, h^{\alpha\beta} = 0, \quad \text{with} \quad h_{\mu\nu;\alpha}{}^{;\alpha} \equiv \Box h_{\mu\nu}.$$

GW lensing: wave optics

BH lenses, historical works, at the formal level :

- Matzner (1968)
- Peters (1976)
- Chrzanowski *et al.* (1976)
- De Logi, Kovacs (1977)
- Futterman *et al.* (1988)

...

More recently: Dolan (2018)

GW lensing: wave optics

Reference work for phenomenology :

Wave effects in gravitational lensing of gravitational waves from chirping binaries

Ryuichi Takahashi (Kyoto U.), Takashi Nakamura (Kyoto U.)
May, 2003

28 pages
Published in: *Astrophys.J.* 595 (2003) 1039-1051
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Scalar wave optics

<u>Usual process</u> : **scalar** wave Ansatz $$h_{\mu\nu} = \phi \cdot e_{\mu\nu}$$

Takahashi, Nakamura (2003), ...

Assume known behaviour for $e_{\mu\nu}$ (parallel transport)

Scalar wave optics

Usual process : **scalar** wave Ansatz $$h_{\mu\nu} = \phi \cdot e_{\mu\nu}$$

Takahashi, Nakamura (2003), ...

Assume known behaviour for $e_{\mu\nu}$ (parallel transport)

Specify the lens background $\bar{g}_{\mu\nu}$ (in our case : Schwarzschild)

Decompose ϕ in multipoles : $$\phi = e^{-i\omega t} \sum_{\ell} \frac{u_\ell(r)}{r\,(1-2M/r)^{1/2}}\, P_\ell(\cos(\theta))$$

Scalar wave optics

Differential equation is

$$\left[\frac{\mathrm{d}^2}{\mathrm{d}r^2} + \omega^2 + \frac{4M\omega^2}{r} - \frac{\ell(\ell+1)}{r^2} + \frac{12M^2\omega^2}{r^2} + \mathcal{O}(r^{-3})\right] u_\ell = 0$$

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Schrödinger in 1/r Coulomb potential (Rutherford) : $\Psi = e^{-i\omega t}\sum_\ell \frac{u_\ell}{r} P_\ell$

$$\left[\frac{\mathrm{d}^2}{\mathrm{d}r^2} + \omega^2 - \frac{2\omega\gamma}{r} - \frac{\ell(\ell+1)}{r^2}\right] u_\ell = 0 \qquad , \ \gamma \propto \text{charges}$$

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Scalar wave optics

Differential equation is

deep wave optics

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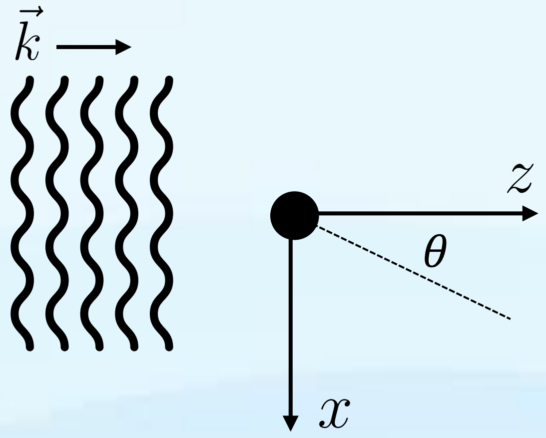
Insights from quantum mechanics

QM problem has exact solution:

$$\Psi \propto e^{-i\omega t} e^{ikz} {}_1F_1[-i\gamma, 1; ikr(1 - \cos(\theta))]$$



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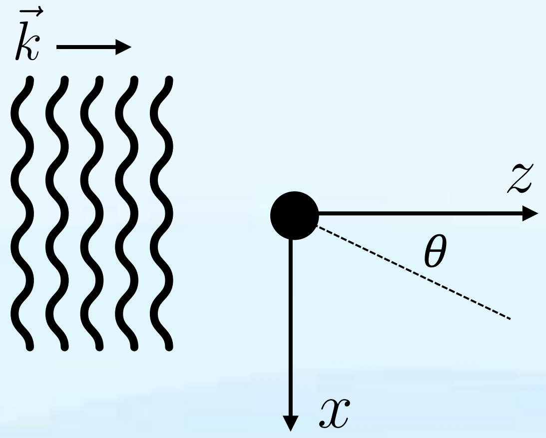

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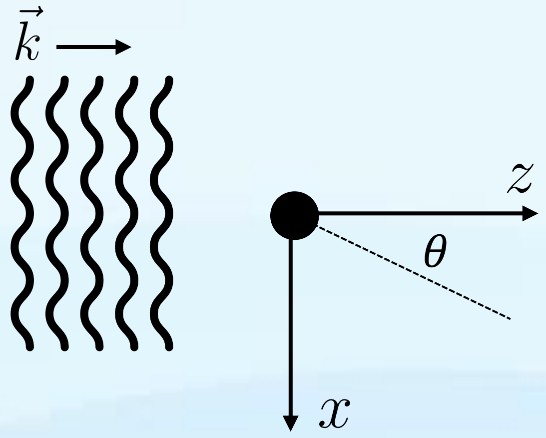
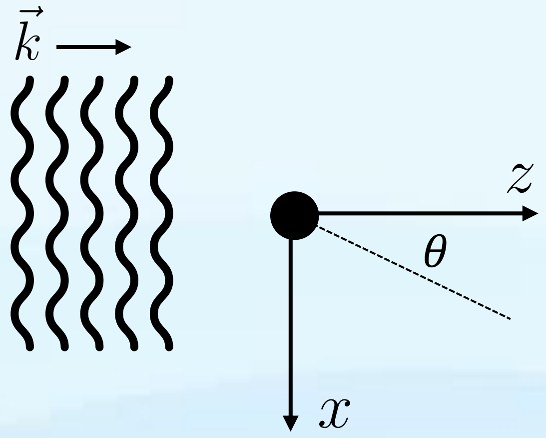


Solution decouples for $kr(1-\cos(\theta)) \gg 1$:

$$e^{i\omega t}\Psi \sim e^{ikz+i\gamma\ln(kr(1-\cos\theta))} - \frac{\tilde{\gamma}}{1-\cos(\theta)} \frac{e^{ikr-i\gamma\ln(kr(1-\cos\theta))}}{kr}$$

Insights from quantum mechanics



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Plane wave

Spherical wave

Insights from quantum mechanics

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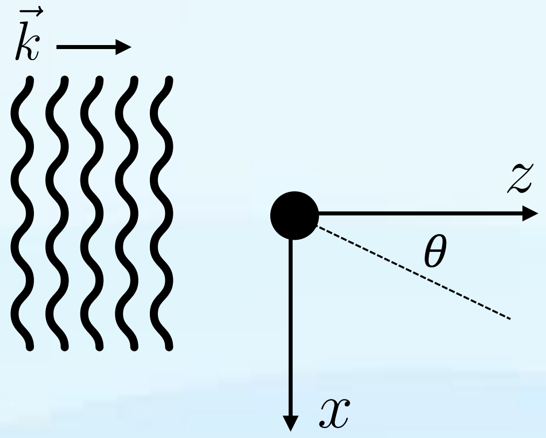

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Plane wave

" amplitude "

Spherical wave

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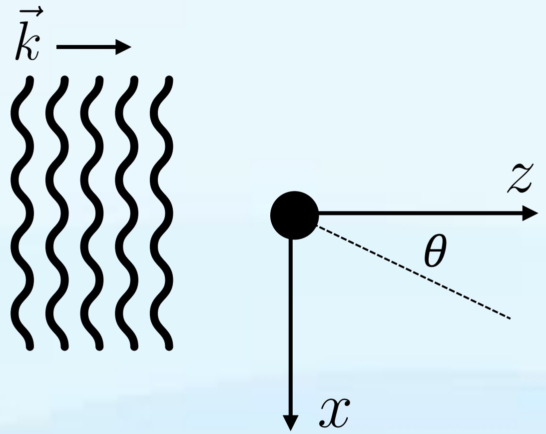

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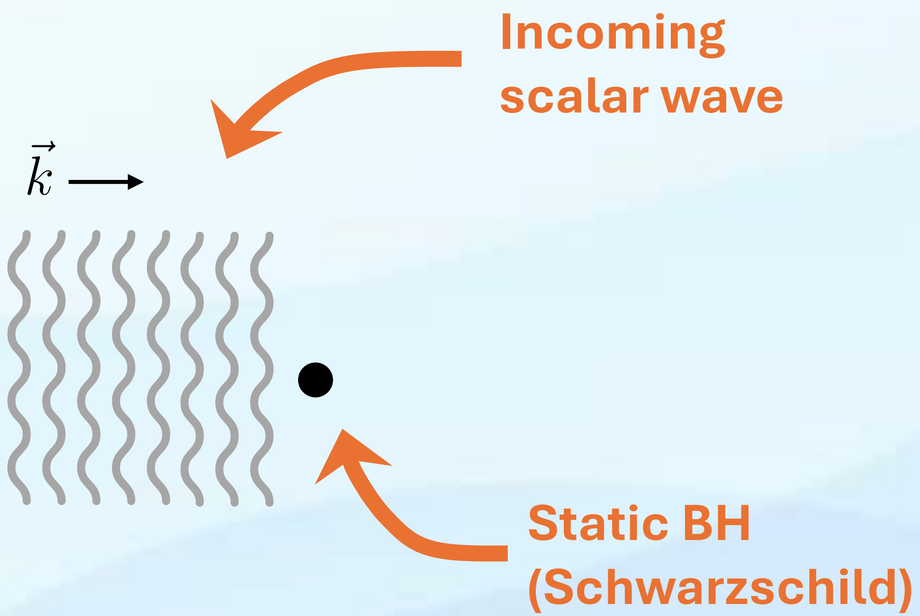
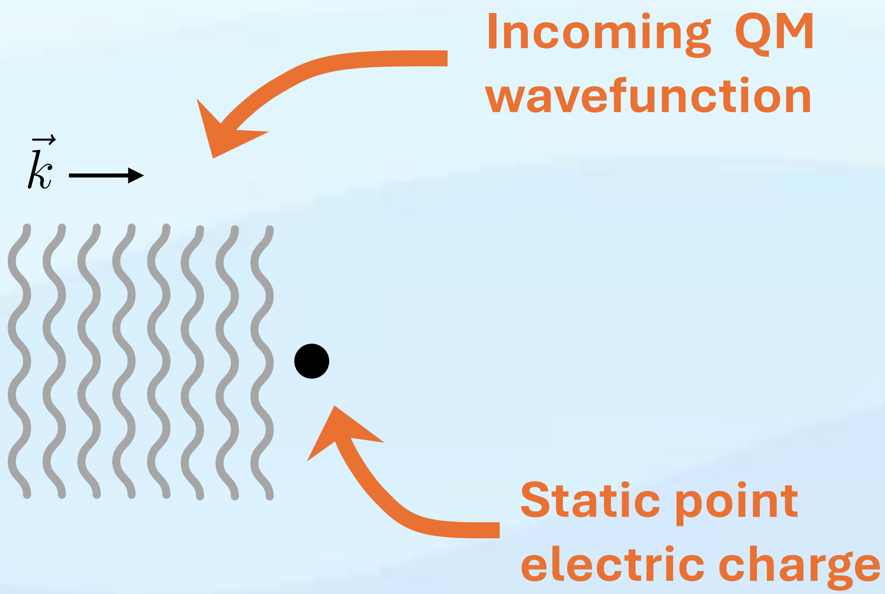
Phase corrections

Insights from quantum mechanics

GR scalar wave, NR simulation



QM, analytical solution



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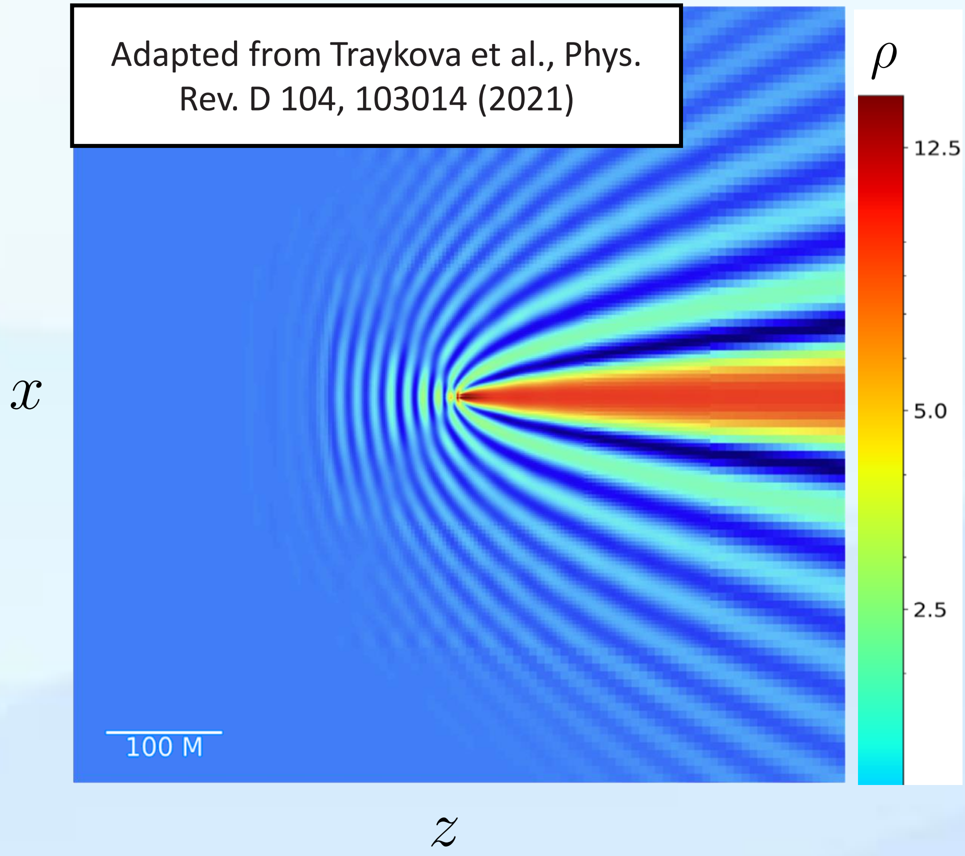

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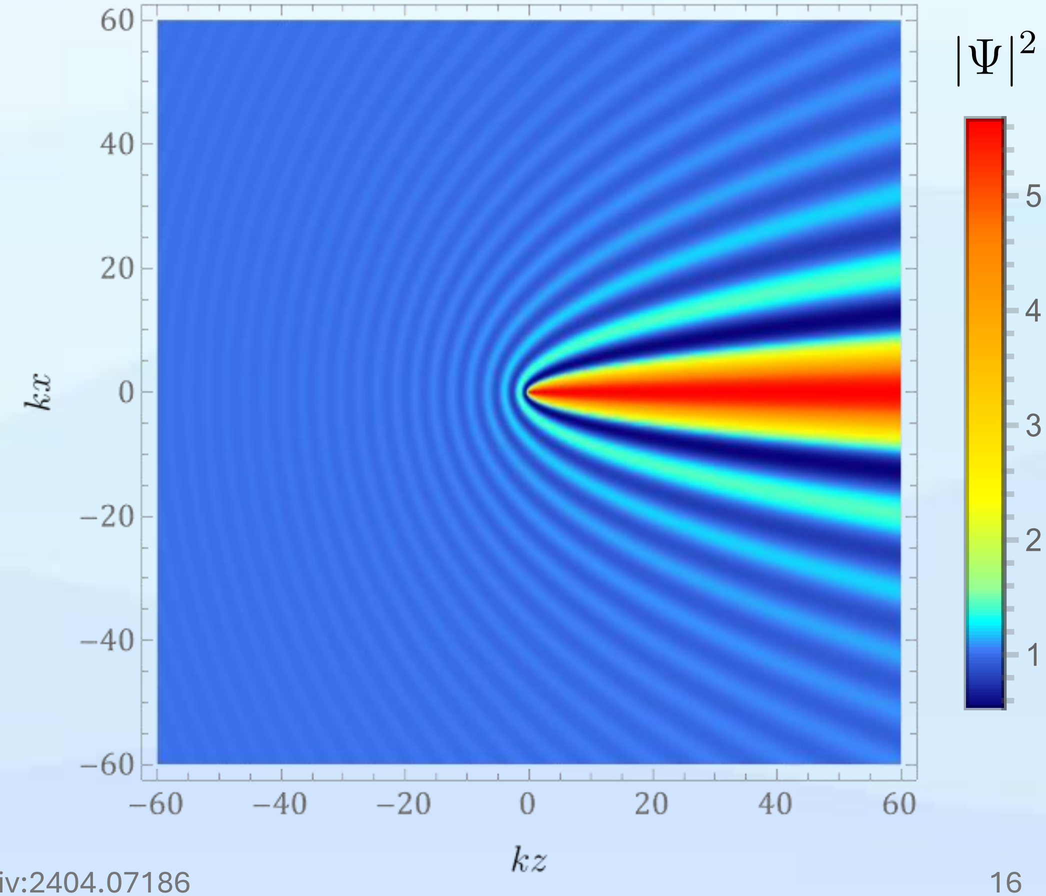

Tensorial wave optics : our work

Start again with : $$\mathrm{d}s^2 = (\bar{g}_{\mu\nu} + h_{\mu\nu})\mathrm{d}x^\mu \mathrm{d}x^\nu$$

... but **avoid** extra assumption $$h_{\mu\nu} = \phi \cdot e_{\mu\nu}$$

Tensorial wave optics : our work

Start again with : $$\mathrm{d}s^2 = (\bar{g}_{\mu\nu} + h_{\mu\nu})\mathrm{d}x^\mu \mathrm{d}x^\nu$$

... but **avoid** extra assumption $h_{\mu\nu} = \phi \cdot e_{\mu\nu}$

Rather treat $h_{\mu\nu}$ with tools of black hole perturbation theory (BHPT),

to keep track of the full polarisation structure

Tensorial wave optics : BHPT

Project $h_{\mu\nu}$ on basis functions on the sphere with even (Y) and odd (X) parity :

$$h_{rr} = \sum_{\ell m} h_{rr}^{\ell m} \, Y^{\ell m},$$ (radial)

$$h_{rA} = \sum_{\ell m} h_{r}^{\ell m} \, X_A^{\ell m} + j_r^{\ell m} \, Y_A^{\ell m}, \quad A = \theta, \phi,$$ (radial/angular)

$$h_{AB} = \sum_{\ell m} h_2^{\ell m} X_{AB}^{\ell m} + r^2 G^{\ell m} Y_{AB}^{\ell m} + r^2 K^{\ell m} \Omega_{AB} Y^{\ell m}, \quad A, B = \theta, \phi,$$ (angular)

Tensorial wave optics : BHPT

From metric multipoles, define two **gauge invariant** master functions:

$$\Psi^{\ell m}_{\text{odd}} = \frac{2r}{(\ell-1)(\ell+2)}\left(\frac{\partial}{\partial r}\hat{h}^{\ell m}_t - \frac{\partial}{\partial t}\hat{h}^{\ell m}_r - \frac{2}{r}\hat{h}^{\ell m}_t\right)$$

$$r^{-1}\Psi^{\ell m}_{\text{even}} \propto \hat{K}^{\ell m} + \frac{2(1-2M/r)}{(\ell-1)(\ell+2)+6M/r}\left((1-2M/r)\hat{h}^{\ell m}_{rr} - r\frac{\partial}{\partial r}\hat{K}^{\ell m}\right)$$

Martel, Poisson. *Physical Review. D* 71.10 (2005)

Tensorial wave optics : BHPT

$\Psi_\bullet^{\ell m}$ obey Zerilli & Regge-Wheeler equations, $\quad \bullet = \text{even, odd}$

$$\frac{\mathrm{d}^2\Psi_\bullet}{\mathrm{d}r_*^2} + \left(\omega^2 - V_\bullet\right)\Psi_\bullet = 0, \qquad \text{with } r_*(r) = r - 2M\ln\left(\frac{r}{2M} - 1\right)$$

Schrödinger-like, for given potentials $V_\bullet(\ell, r, M)$

Poisson, Sasaki. *Physical Review D* 51.10 (1995)

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Schrödinger-like, for given potentials $V_\bullet(\ell, r, M)$

For the scattering problem :

Asymptotic solutions **for $\boldsymbol{\omega M \ll 1}$** are known, expect $\Psi_\bullet^{\ell m} \sim \Psi_\bullet^{\text{plane}} + \Psi_\bullet^{\text{sph}}$

Poisson, Sasaki. *Physical Review D* 51.10 (1995)

Tensorial wave optics : BHPT

Assume initial

$$h_{\mu\nu}^{\text{source}}$$

(absence of lens)

Derive corresponding

$$\Psi_\bullet^{\text{source}}$$

Solve differential equ. such that

$$\Psi_\bullet^{\ell m} \sim \Psi_\bullet^{\text{source}} + \Psi_\bullet^{\text{lensed}}$$

Derive final

$$h_{\mu\nu}$$

(incl. lensing)

Tensorial wave optics : BHPT

« Observables »

Assume initial

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Tensorial wave optics : BHPT

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$\longrightarrow$

Derive corresponding

$$\Psi_{\bullet}^{\text{source}}$$

$\downarrow$

Solve differential equ. such that

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$\longleftarrow$

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Derive final

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Tensorial wave optics : BHPT

Assume initial

$$h_{\mu\nu}^{\text{source}}$$

(absence of lens)

0PN is fine here

→

Derive corresponding

$$\Psi_{\bullet}^{\text{source}}$$

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Solve differential equ. such that

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Tensorial wave optics : BHPT

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Derive final

$$h_{\mu\nu}$$

(incl. lensing)

Tensorial wave optics : BHPT

Assume initial

$$h_{\mu\nu}^{\text{source}}$$

(absence of lens)

$\xrightarrow{\textbf{Project on sphere}}$

Derive corresponding

$$\Psi_{\bullet}^{\text{source}}$$

$\downarrow$

Solve differential equ. such that

$$\Psi_{\bullet}^{\ell m} \sim \Psi_{\bullet}^{\text{source}} + \Psi_{\bullet}^{\text{lensed}}$$

$\leftarrow$

Derive final

$$h_{\mu\nu}$$

(incl. lensing)

Tensorial wave optics : BHPT

Assume initial

$h_{\mu\nu}^{\text{source}}$

(absence of lens)

Derive corresponding

$\Psi_{\bullet}^{\text{source}}$

Solve differential equ. such that

$\Psi_{\bullet}^{\ell m} \sim \Psi_{\bullet}^{\text{source}} + \Psi_{\bullet}^{\text{lensed}}$

Derive final

$h_{\mu\nu}$

(incl. lensing)

Tensorial wave optics : BHPT

Assume initial

$h_{\mu\nu}^{\text{source}}$

(absence of lens)

Derive corresponding

$\Psi_{\bullet}^{\text{source}}$

Using scattering phase shifts

Solve differential equ. such that

$\Psi_{\bullet}^{\ell m} \sim \Psi_{\bullet}^{\text{source}} + \Psi_{\bullet}^{\text{lensed}}$

Derive final

$h_{\mu\nu}$

(incl. lensing)

Tensorial wave optics : BHPT

Assume initial

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(absence of lens)

Derive corresponding

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Solve differential equ. such that

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(absence of lens)

Derive corresponding

$$\Psi_{\bullet}^{\text{source}}$$

Solve differential equ. such that

$$\Psi_{\bullet}^{\ell m} \sim \Psi_{\bullet}^{\text{source}} + \Psi_{\bullet}^{\text{lensed}}$$

Hidden work: facing divergent series

non standard summation methods

Derive final

$$h_{\mu\nu}$$

(incl. lensing)

Familiar features recovered

Scalar QM problem had closed form solution:

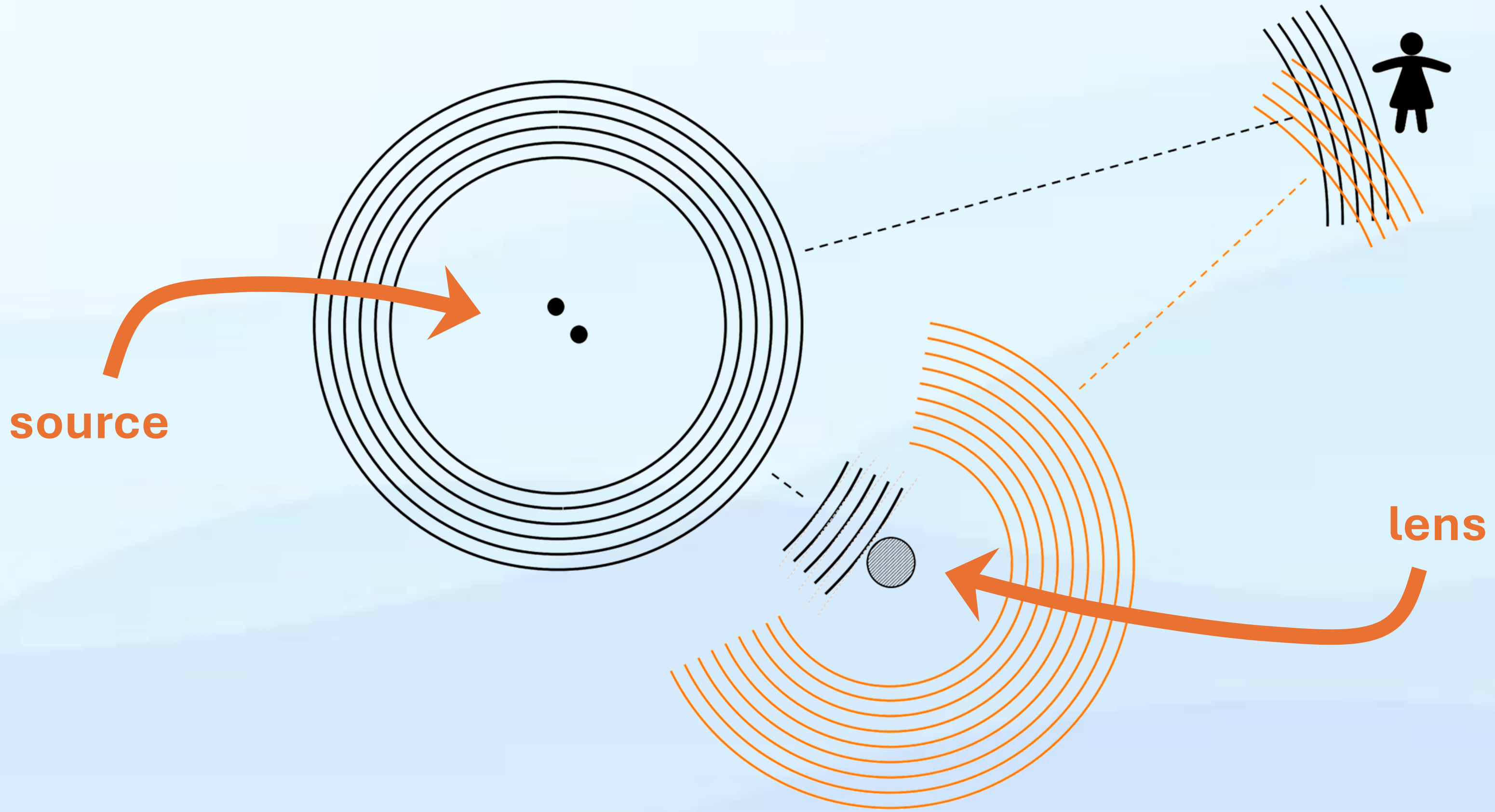
$$\Psi \propto e^{-i\omega t} e^{ikz} {}_1F_1[-i\gamma, 1; ikr(1-\cos(\theta))]$$

Our ***asymptotical*** **GW solution shares asymptotic features with the latter …**

… Extending them in a spin-2 version (full expressions in arXiv:2404.07186)

Tensorial wave optics : interference

Full wave solution is a **superposition** of lensed and original wave



Polarisation

Quantifying the signal polarisation content $\mathcal{V} \in [-1, 1]$:

$$\mathcal{V} \equiv \frac{2\mathrm{Im}[\tilde{h}_+ \tilde{h}_\times^*]}{|\tilde{h}_+|^2 + |\tilde{h}_\times|^2} = V/I \qquad \text{in terms of the Stokes parameters } V, I\,.$$

$$= \frac{|\tilde{h}^{(2)}|^2 - |\tilde{h}^{(-2)}|^2}{|\tilde{h}^{(2)}|^2 + |\tilde{h}^{(-2)}|^2}$$

constant in geometric optics and scalar wave optics

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constant in geometric optics and scalar wave optics

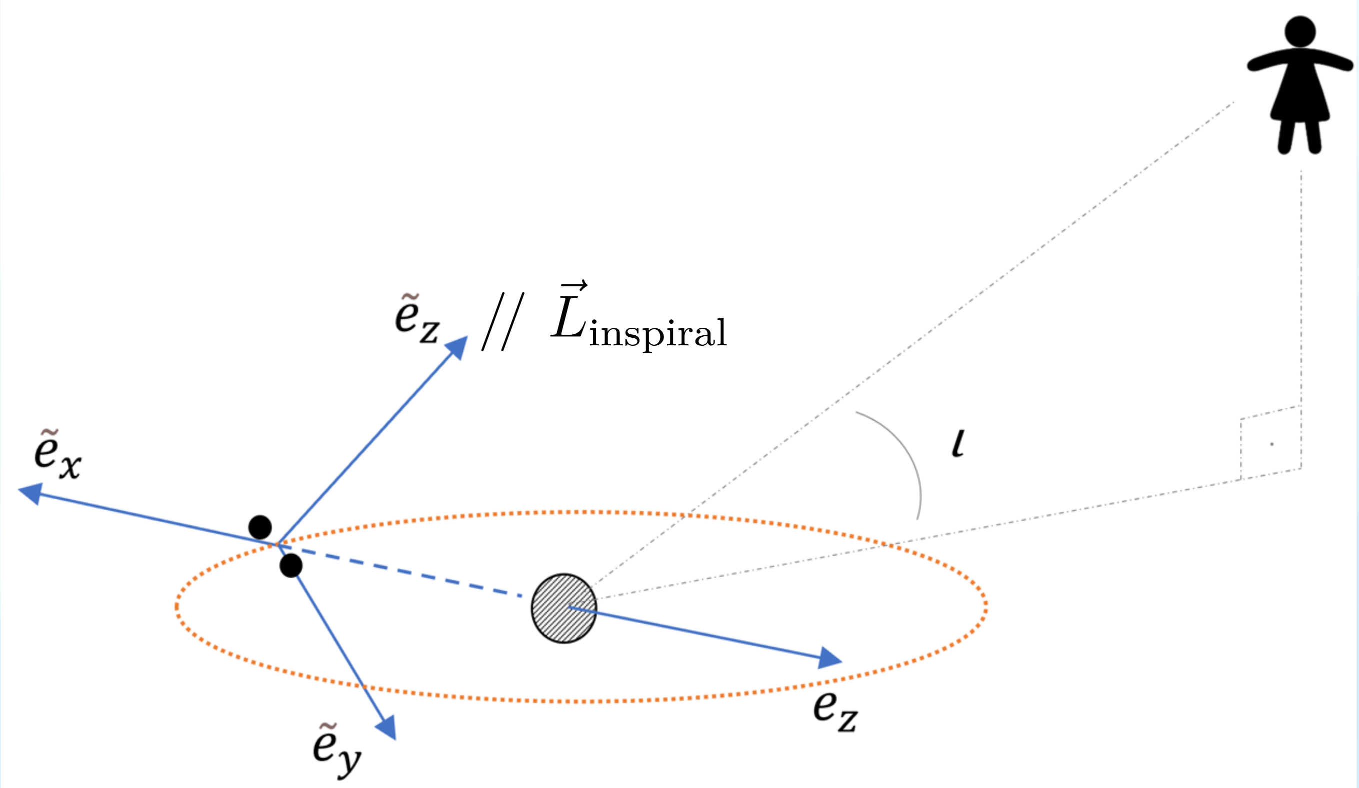
in general **not constant** tensorial wave optics

Wave optics lensing in triple systems: towards a phenomenology

Illustration:

- For simple circular orbits (disk migration traps, alike GW190521 ?)

- Considering a time-varying alignment (*explicitly ignoring other velocity dependent terms, e.g. Doppler*)



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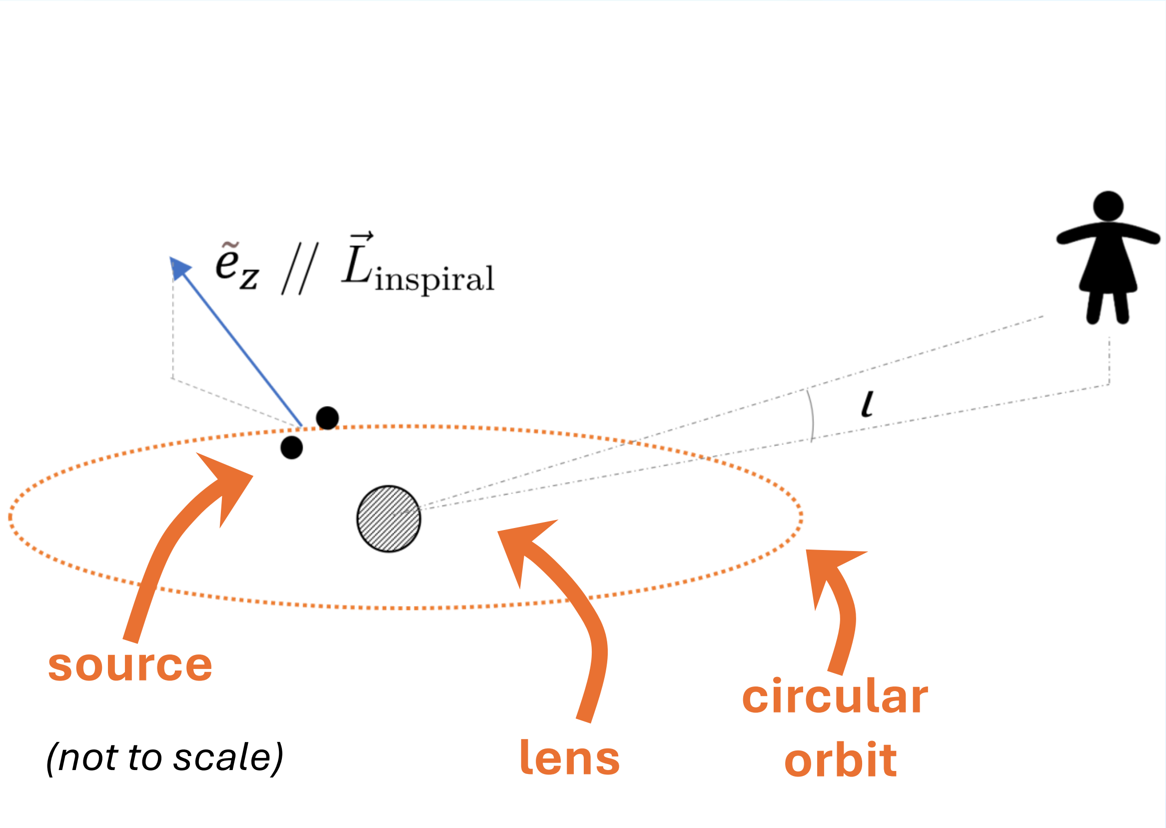
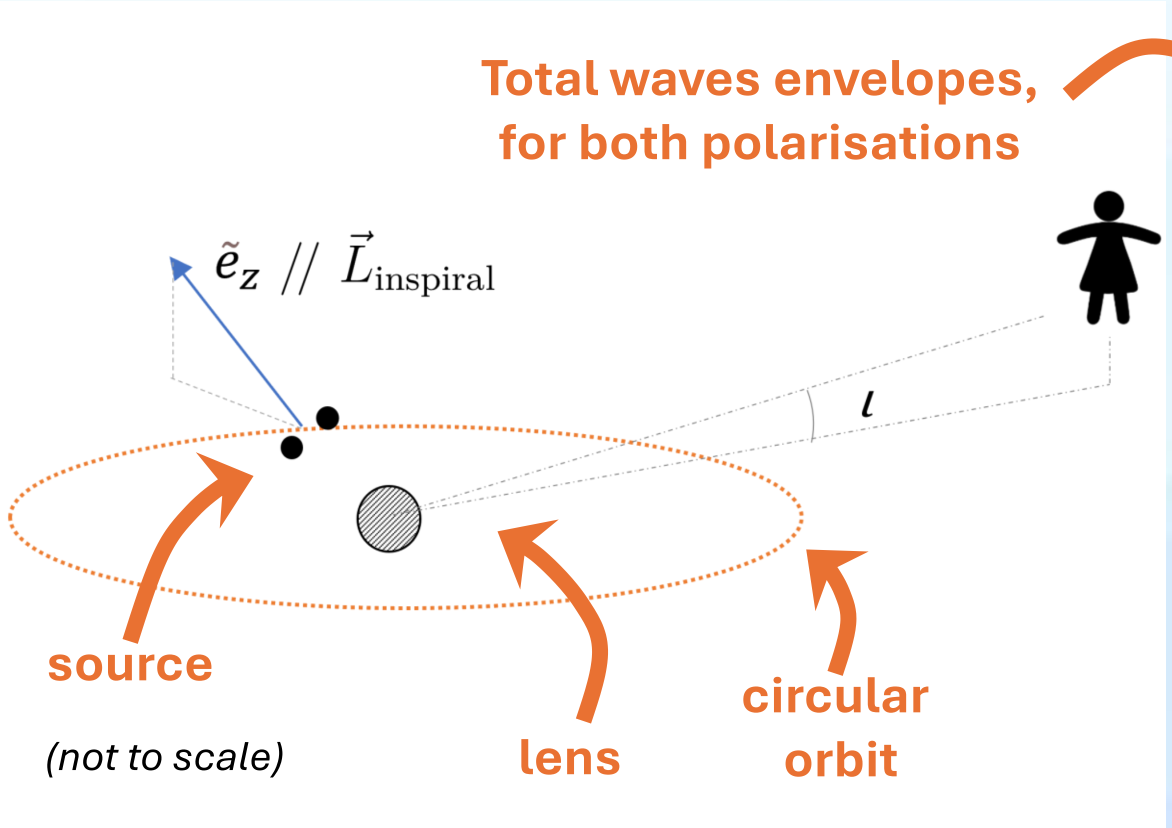
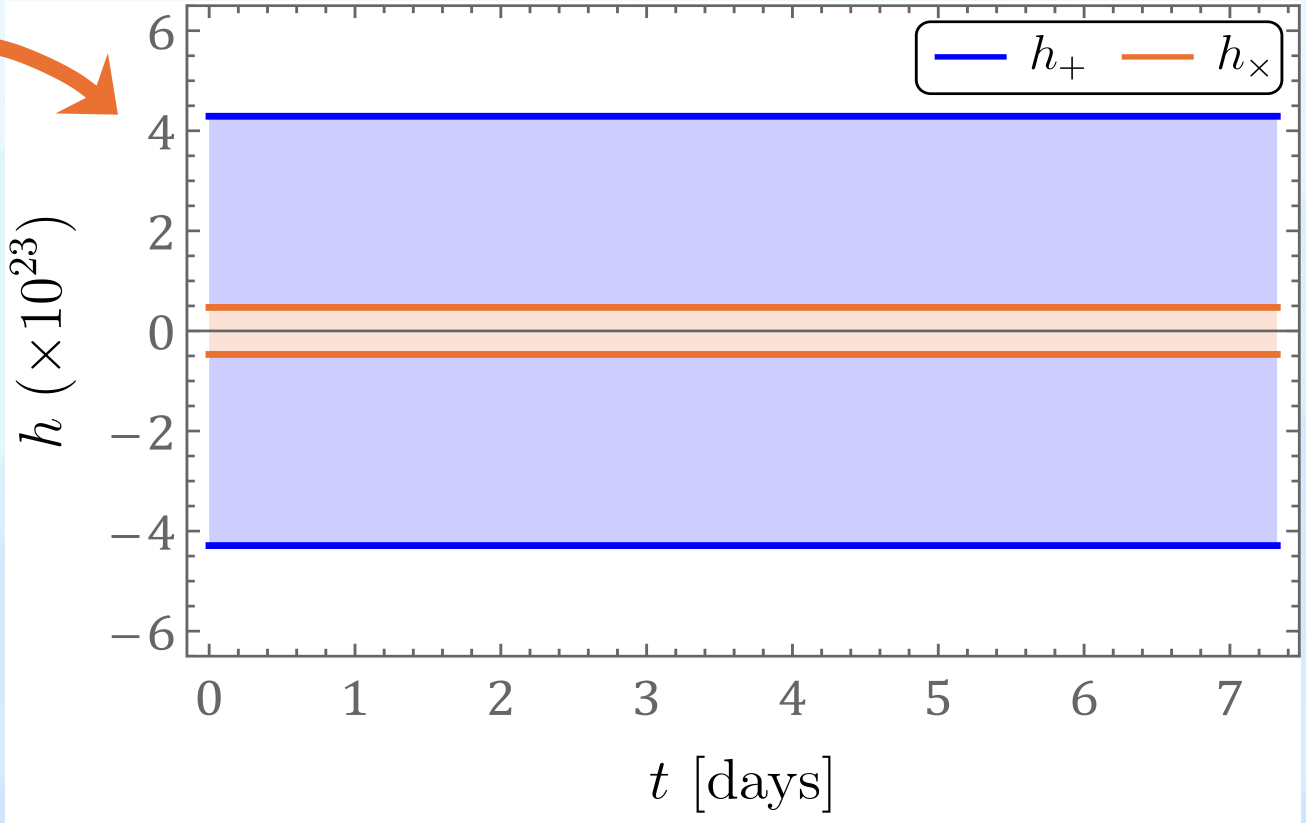


Wave optics lensing in triple systems: towards a phenomenology

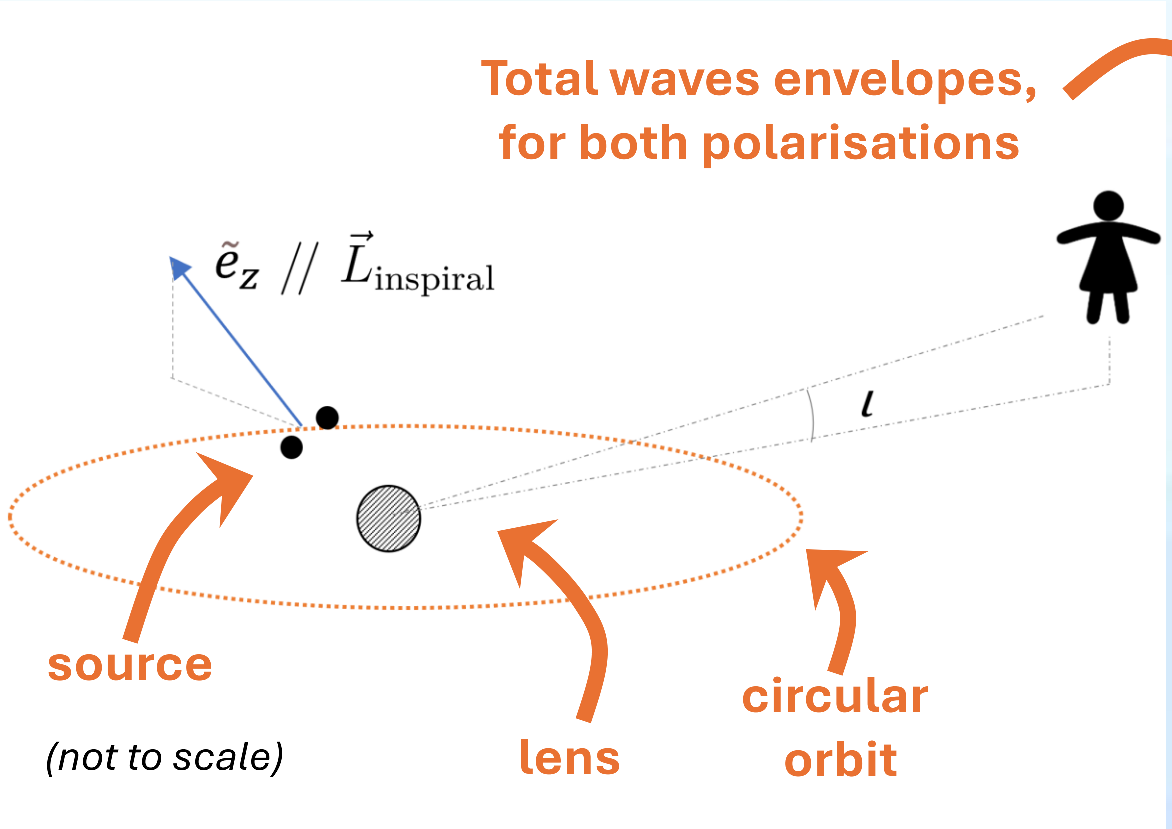
Inspiral phase *without lens*

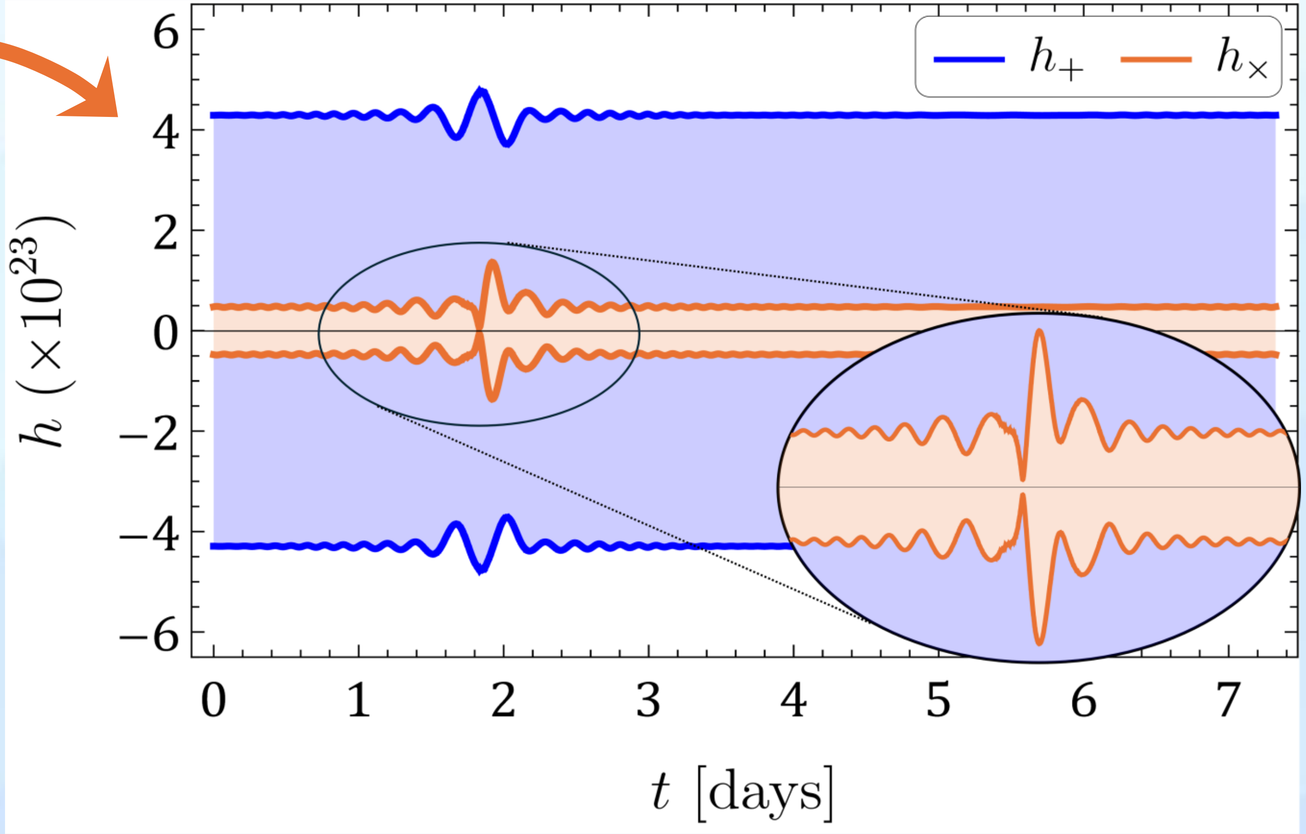




Wave optics lensing in triple systems: towards a phenomenology

Toy GW190521-inspired source, in LISA- « optimal » wave optics





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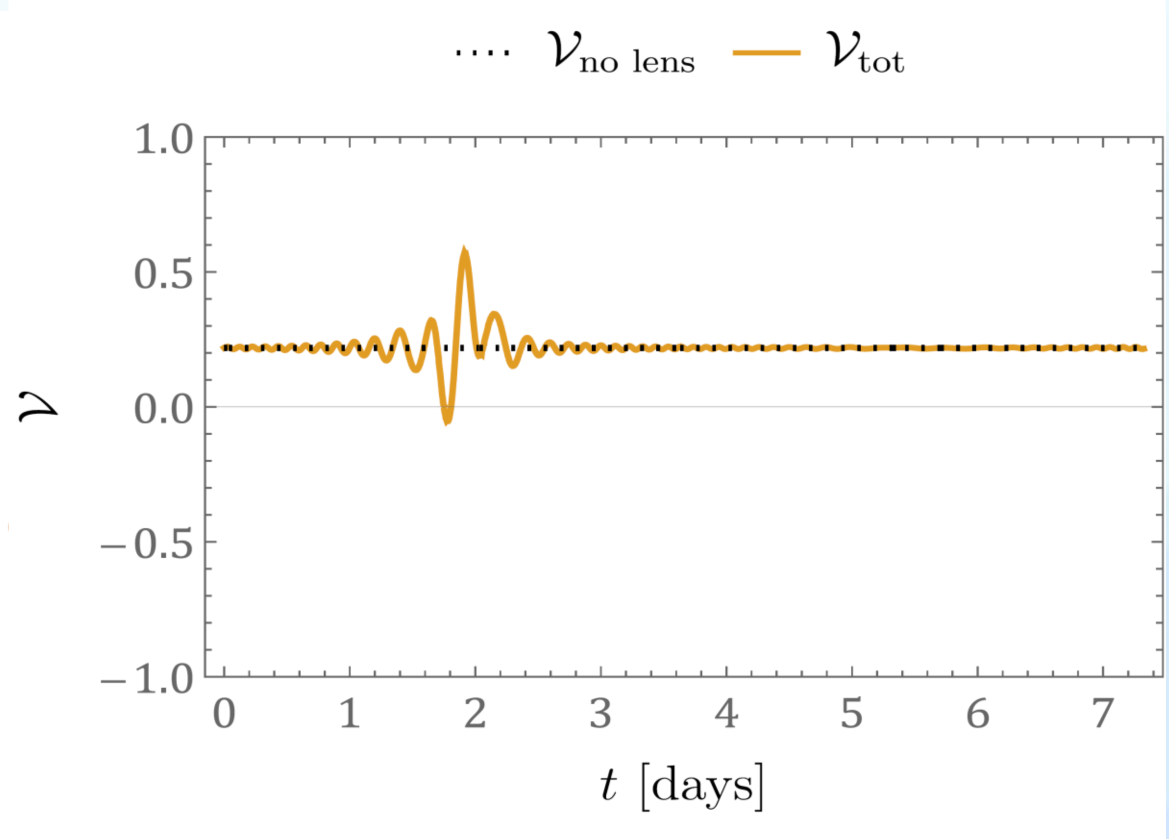
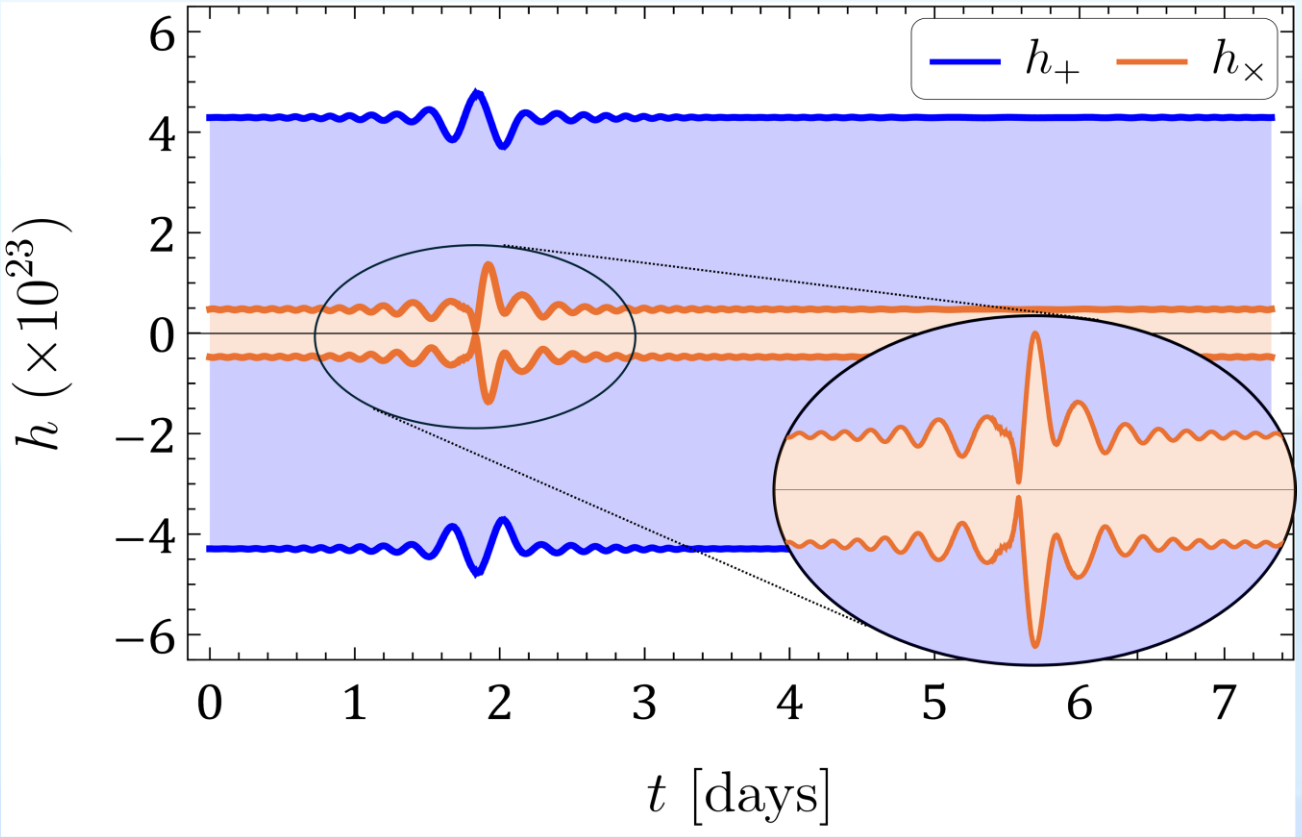
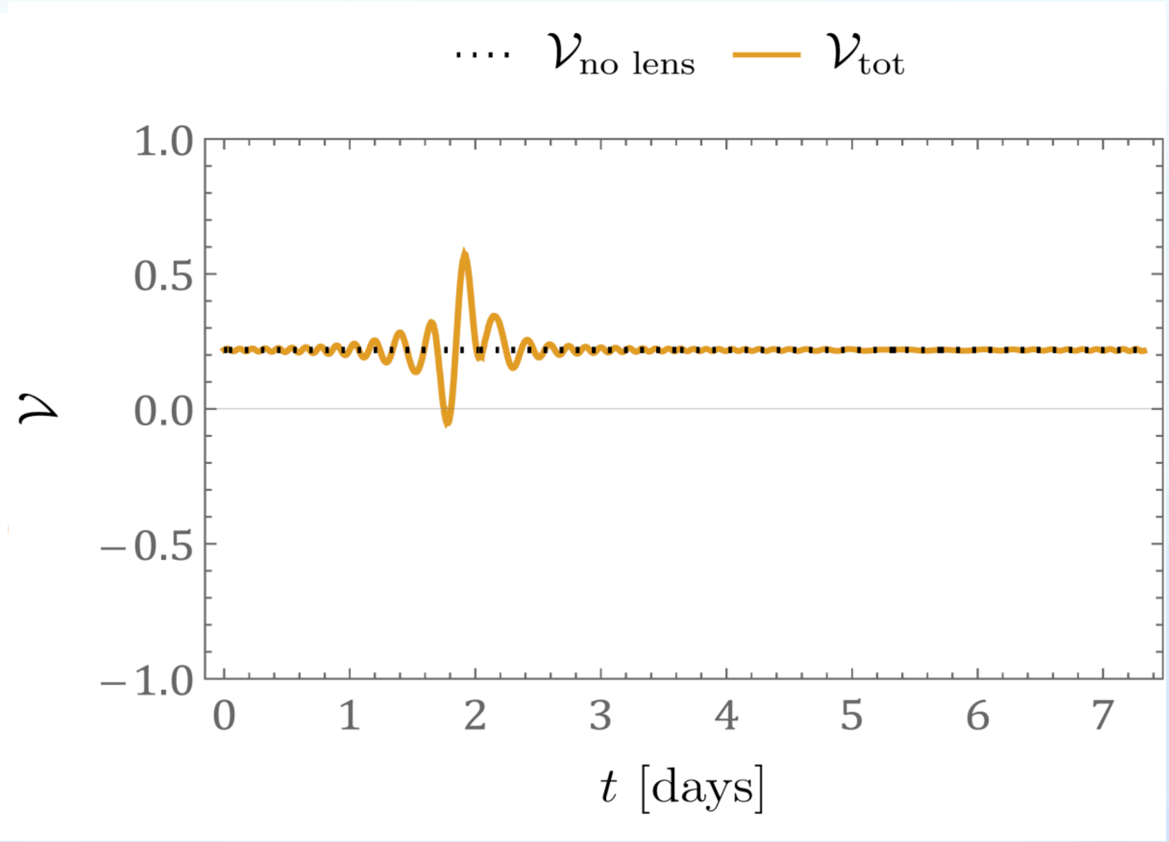
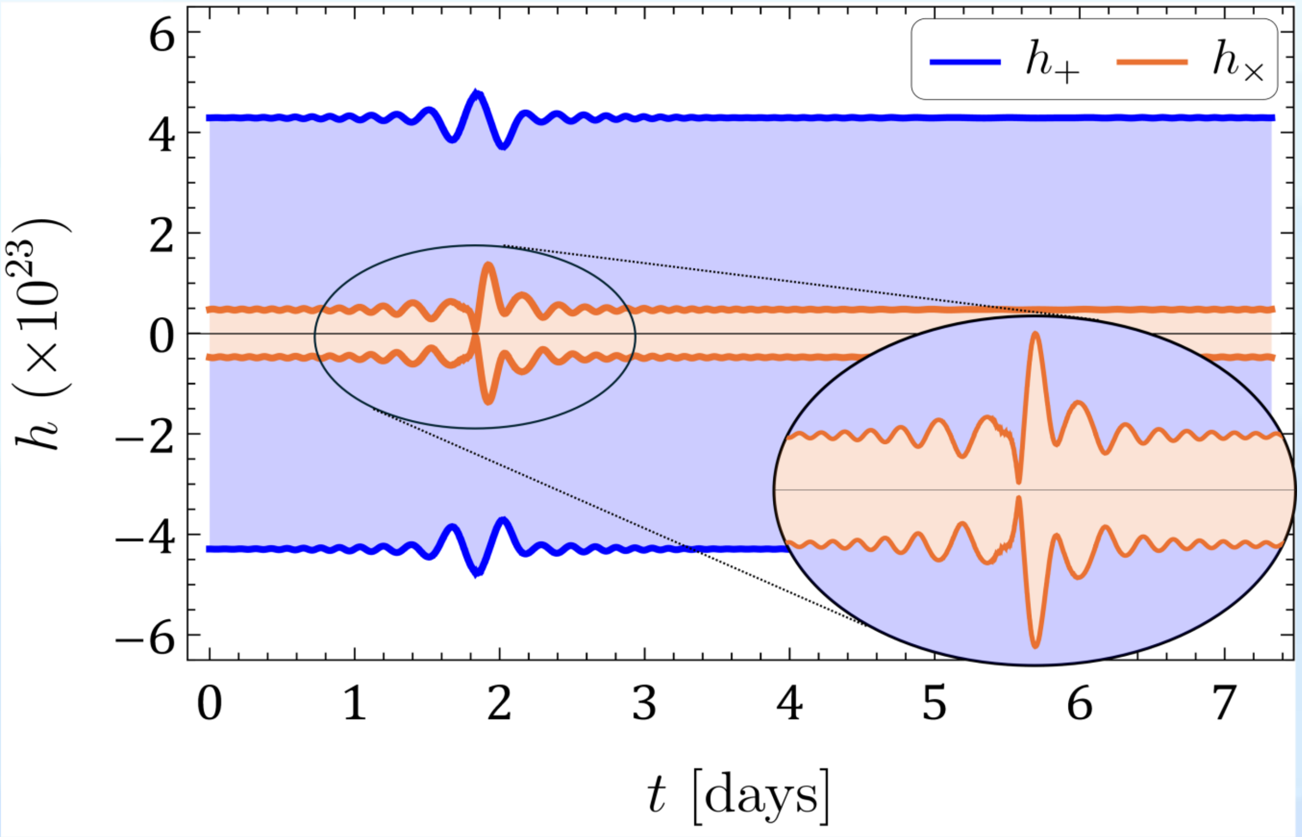




Wave optics lensing in triple systems: towards a phenomenology

LISA **detectable** with SNR > 100 if at $z \sim 0.01$ (closest AGN)





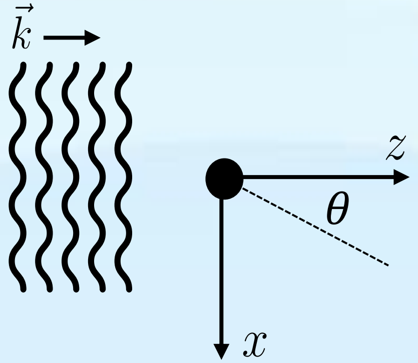
Conclusions

- **Wave optics** is a regime in the range of future observations (LISA)
- **Triple systems** exhibit a rich dynamical lensing phenomenology
- In wave optics, GW lensing is a **fully tensorial** process:
  - Results go beyond a simple scalar amplification factor
  - **The polarisation/helicity structure is not preserved by lensing**

Polarisation

Quantifying the signal polarisation content $\mathcal{V} \in [-1, 1]$:

$$\mathcal{V} \equiv \frac{2\mathrm{Im}[\tilde{h}_+ \tilde{h}_\times^*]}{|\tilde{h}_+|^2 + |\tilde{h}_\times|^2} = V/I$$



Wave optics lensing is **polarisation dependent**, e.g. :

$$\frac{\mathrm{d}\sigma}{\mathrm{d}\Omega} = M^2 \frac{\cos^8\left(\frac{\theta}{2}\right) + \sin^8\left(\frac{\theta}{2}\right)}{\sin^4\left(\frac{\theta}{2}\right)} + 2M^2\sqrt{1 - \mathcal{V}^2_{\text{incident}}}\,\cos^4\left(\frac{\theta}{2}\right)\cos(4\phi)$$

Wave optics lensing in triple systems: towards a phenomenology

- LISA-band system : $\omega = 2\pi f = 2\pi \times 3 \times 10^{-3}\mathrm{Hz}$

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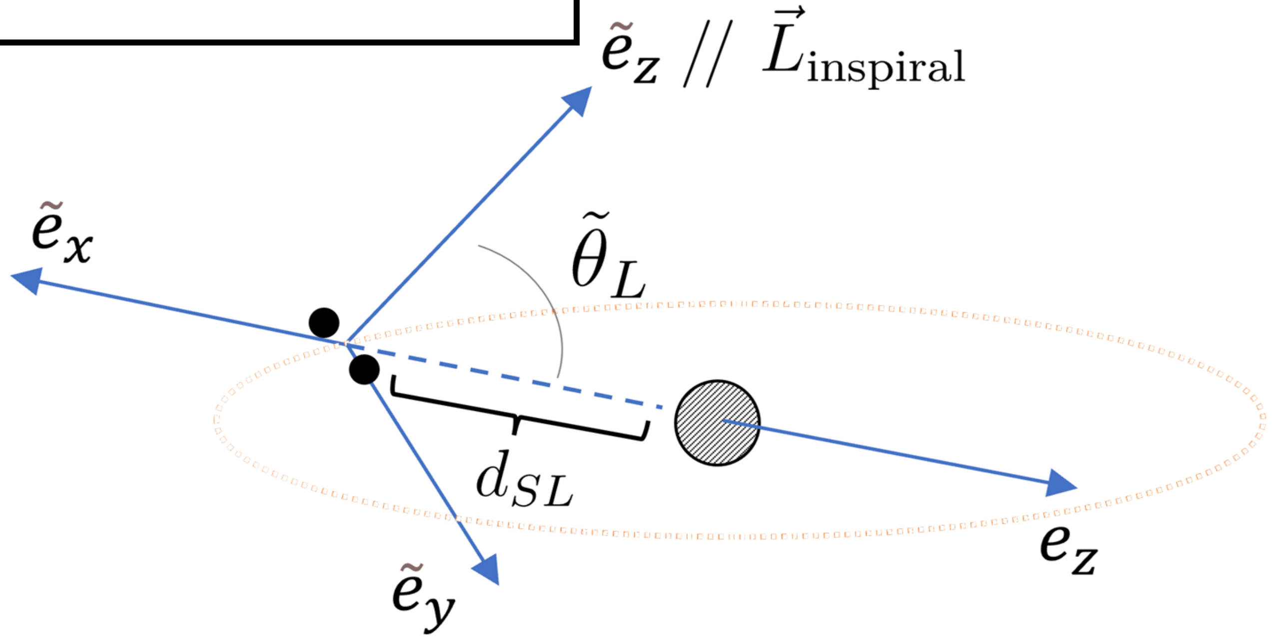
LISA **detectable** with SNR > 100 if at $z \sim 0.01$ (nearest AGNs)

Tensorial wave optics : BHPT

Assume initial

$$h_{\mu\nu}^{\text{source}}$$

(absence of lens)



TT gauge, propagation along e_z :

$$h_{ij}^{\text{source}} = \begin{pmatrix} h_+ & h_\times & 0 \\ h_\times & -h_+ & 0 \\ 0 & 0 & 0 \end{pmatrix}_{ij}$$

$$h_+ = \frac{A_{\text{in}}}{\tilde{r}} \frac{1+\cos^2\tilde{\theta}_L}{2} \cos[\omega(t-\tilde{r}) - 2\tilde{\phi}_L]$$

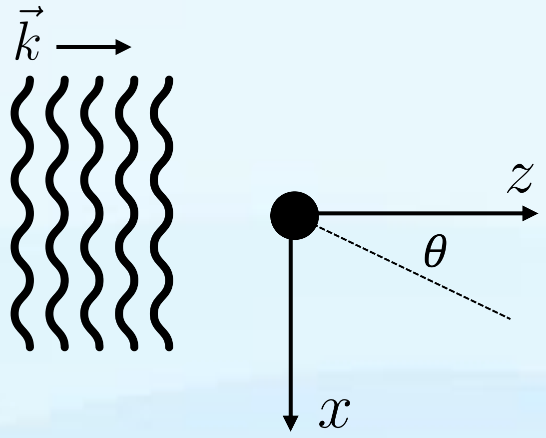
$$h_\times = \frac{A_{\text{in}}}{\tilde{r}} \cos\tilde{\theta}_L \sin[\omega(t-\tilde{r}) - 2\tilde{\phi}_L]$$

(locally plane wave)

Insights from quantum mechanics

QM problem has exact solution:

$$\Psi \propto e^{-i\omega t} e^{ikz} {}_1F_1[-i\gamma, 1; ikr(1-\cos(\theta))]$$



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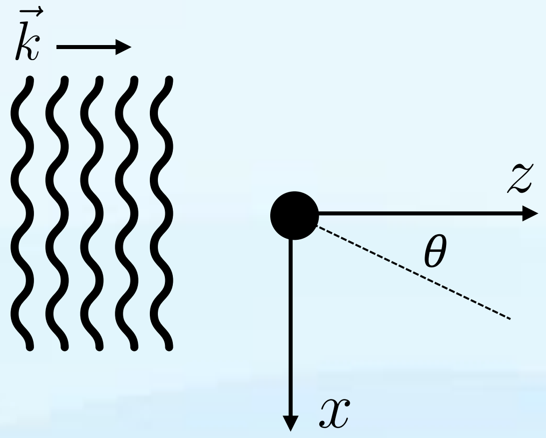

Solution decouples for $kr(1-\cos(\theta)) \gg 1$:

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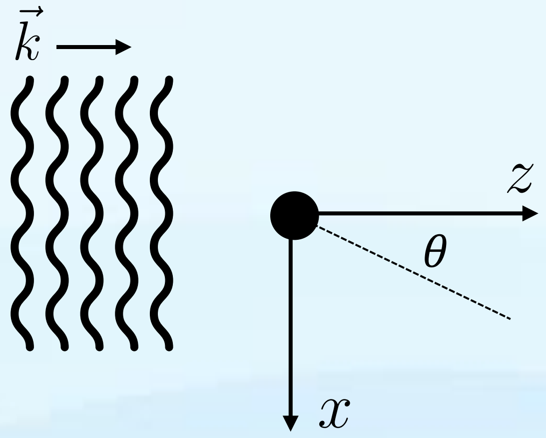

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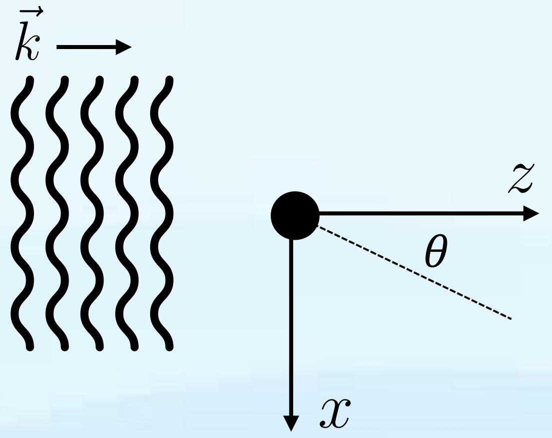

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Plane wave

Spherical wave

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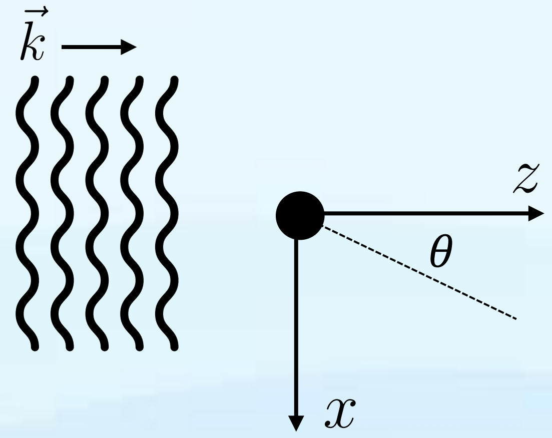

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apparent divergence as $\theta \to 0$

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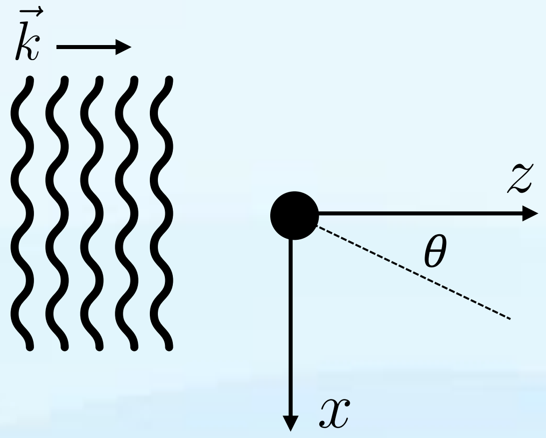

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log phase corrections

Familiar features recovered

Scalar QM problem had solution:

apparent divergence as $\theta \to 0$

$$e^{i\omega t}\Psi \sim e^{ikz+i\gamma\ln kr(1-\cos(\theta))} - \frac{\tilde{\gamma}}{kr(1-\cos(\theta))}e^{ikr-i\gamma\ln kr(1-\cos(\theta))}$$

log phase corrections

Our ***asymptotical*** **GW solution recovers a spin-2 version** of these features (full expressions in arXiv:2404.07186)

Tensorial wave optics : results

Recovering familiar features:

$$h_+ - h_+^{\text{source}} - \simeq \frac{A_{\text{in}}}{r}\frac{2M}{d_{\text{SL}}}\frac{1+\cos^2\theta}{2}\frac{1}{1-\cos\theta}\left(\cos^4\left(\frac{\tilde{\theta}_L}{2}\right)\cos\left(\varphi - 2\phi\right) + \sin^4\left(\frac{\tilde{\theta}_L}{2}\right)\cos\left(\varphi + 2\phi\right)\right)$$

$$h_\times - h_\times^{\text{source}} - \simeq \frac{A_{\text{in}}}{r}\frac{2M}{d_{\text{SL}}}\cos\theta\frac{1}{1-\cos\theta}\left(\cos^4\left(\frac{\tilde{\theta}_L}{2}\right)\sin\left(\varphi - 2\phi\right) - \sin^4\left(\frac{\tilde{\theta}_L}{2}\right)\sin\left(\varphi + 2\phi\right)\right)$$

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TT gauge projection
of a quadrupole
(cf source)

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lensing features

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Natural *expected* validity range : $kr(1-\cos(\theta)) \gg 1$

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log phase corrections

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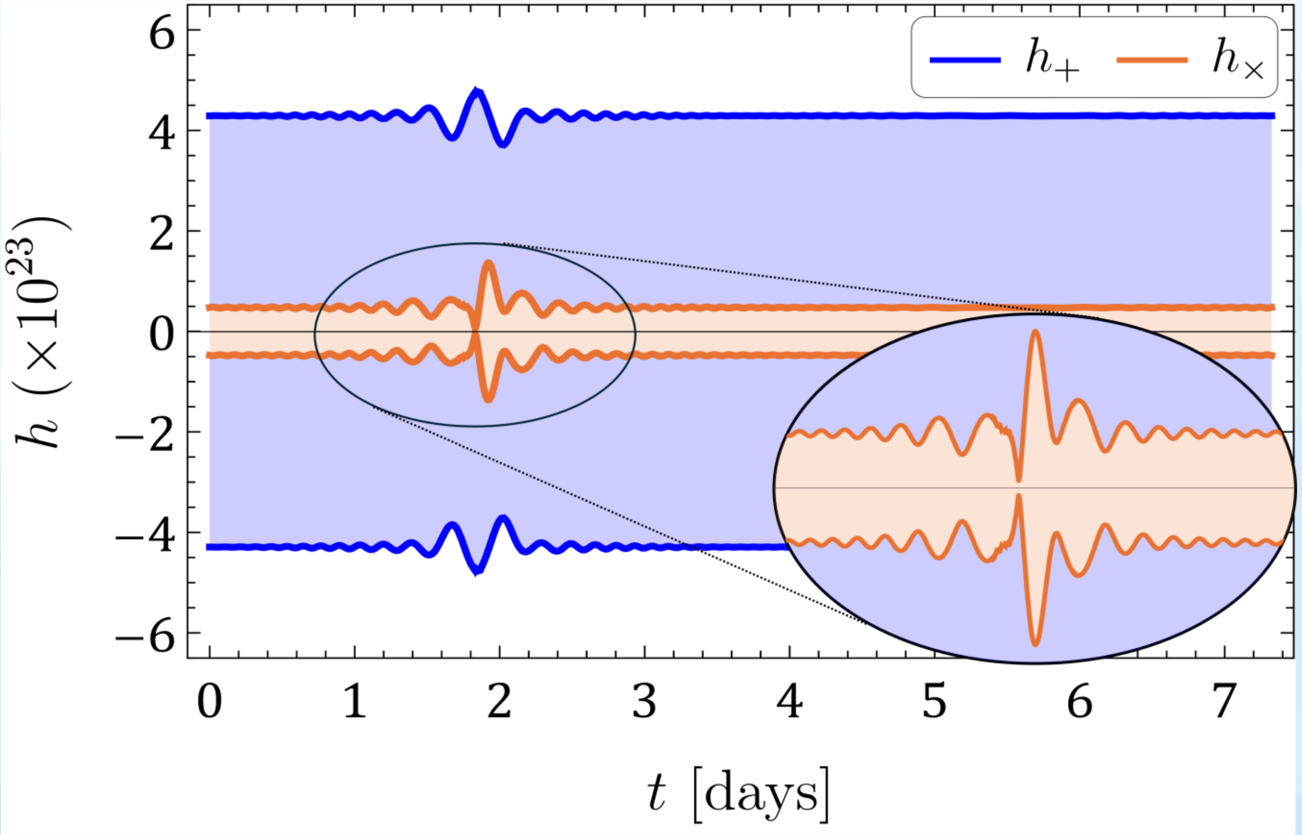
Toy GW190521-inspired source, in LISA- « optimal » wave optics

Total waves envelopes, for both polarisations



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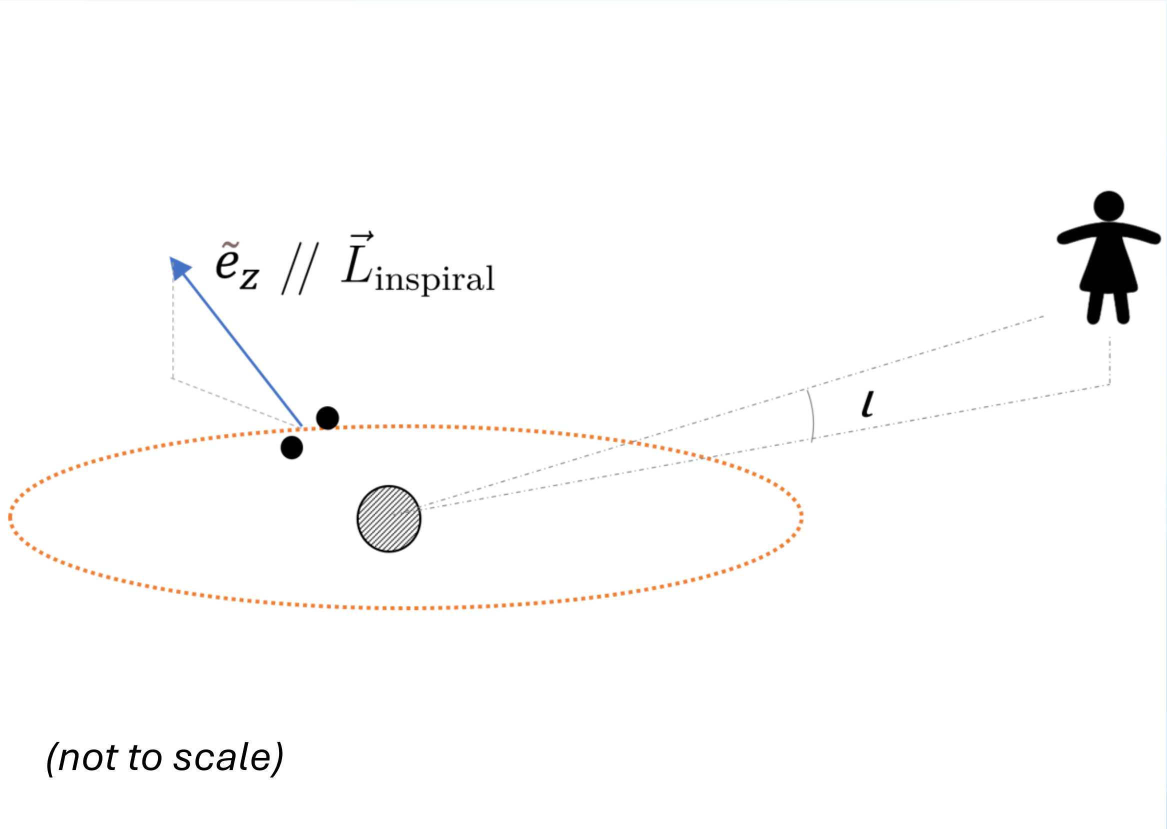

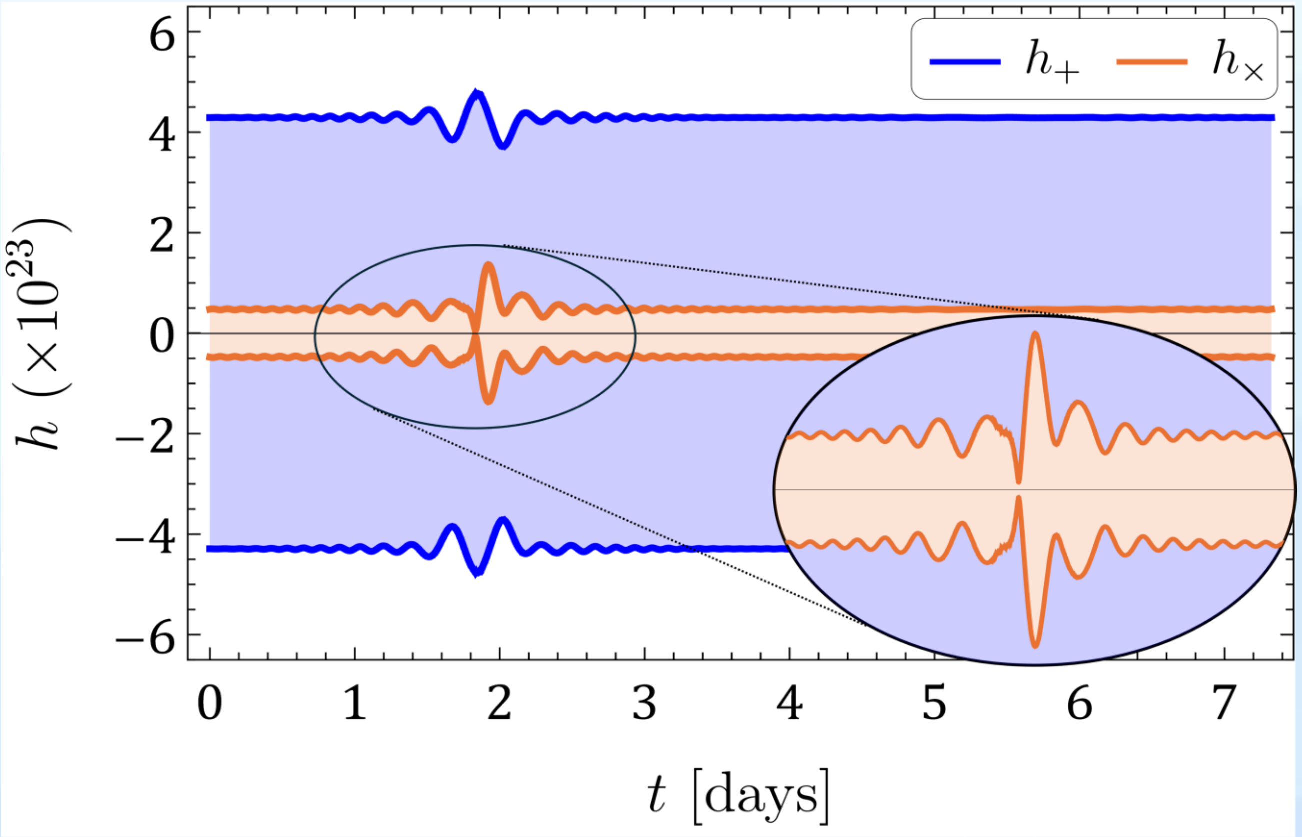

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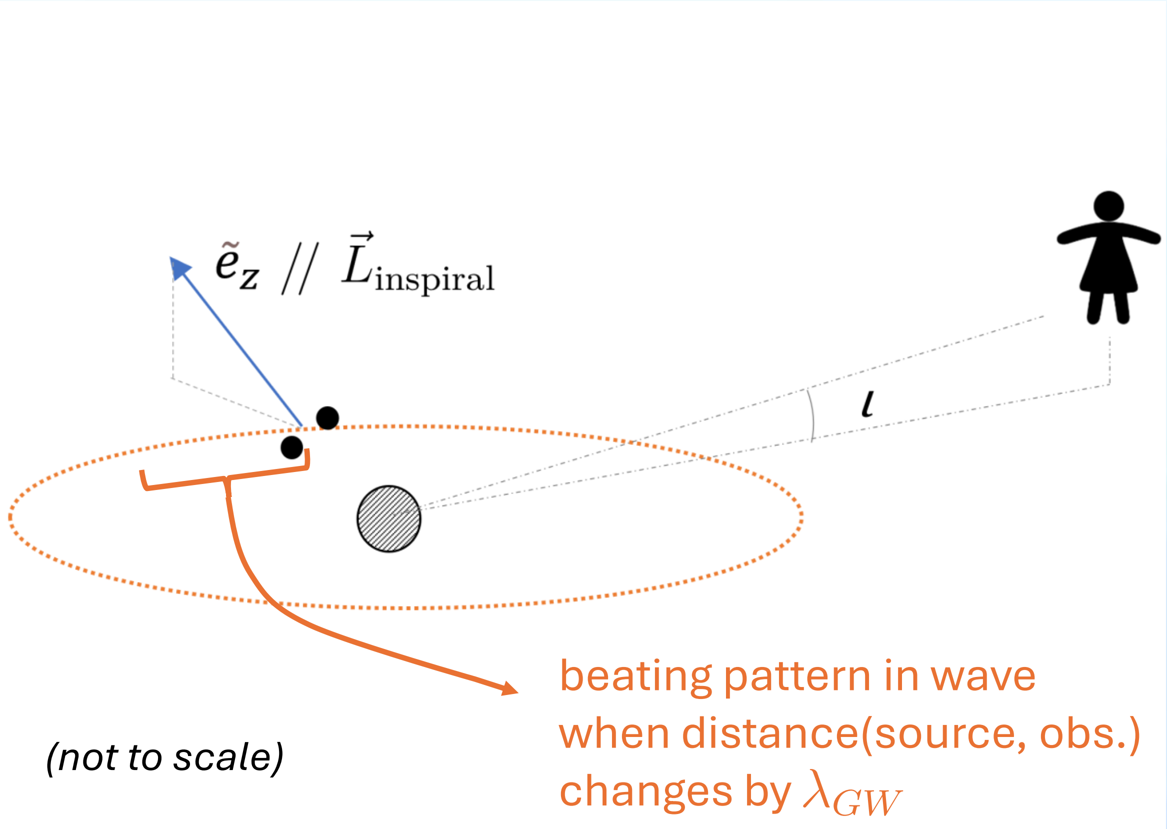

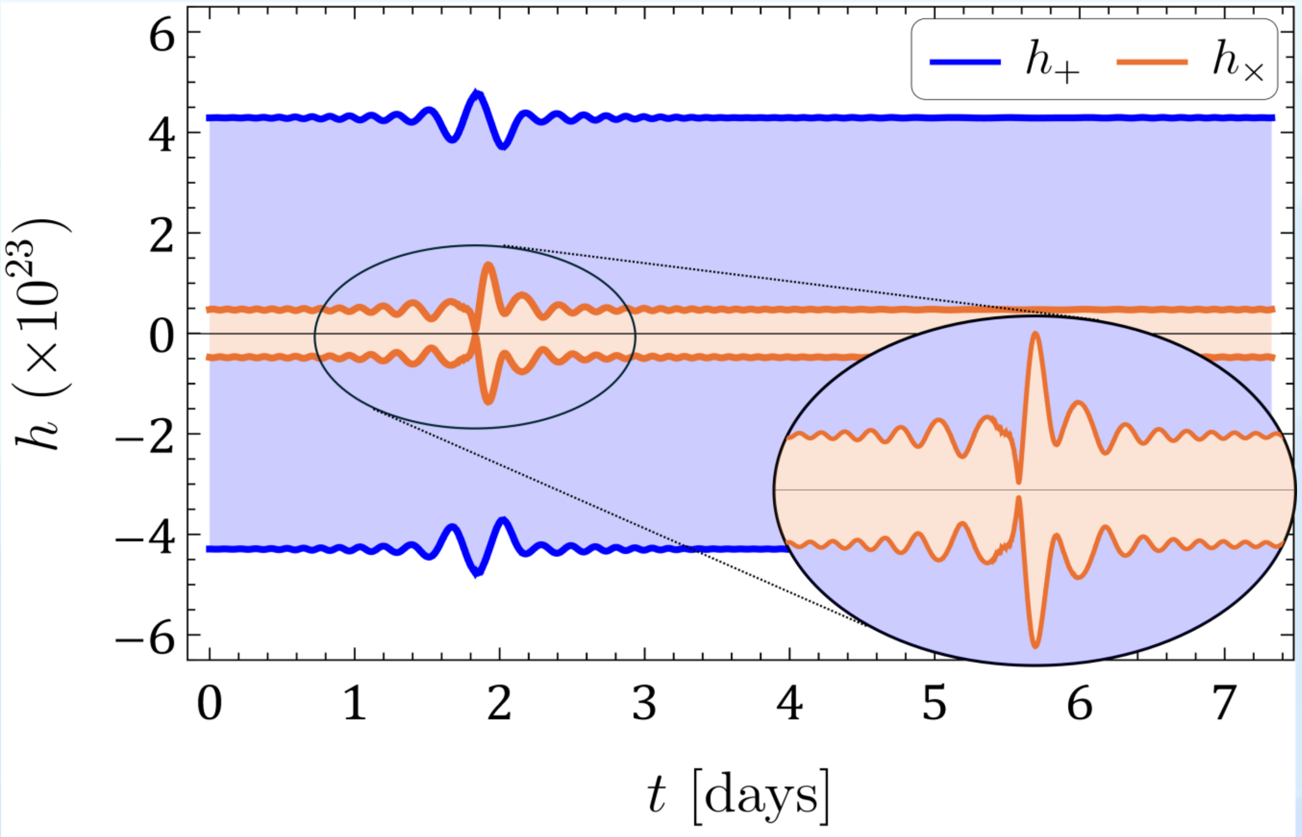

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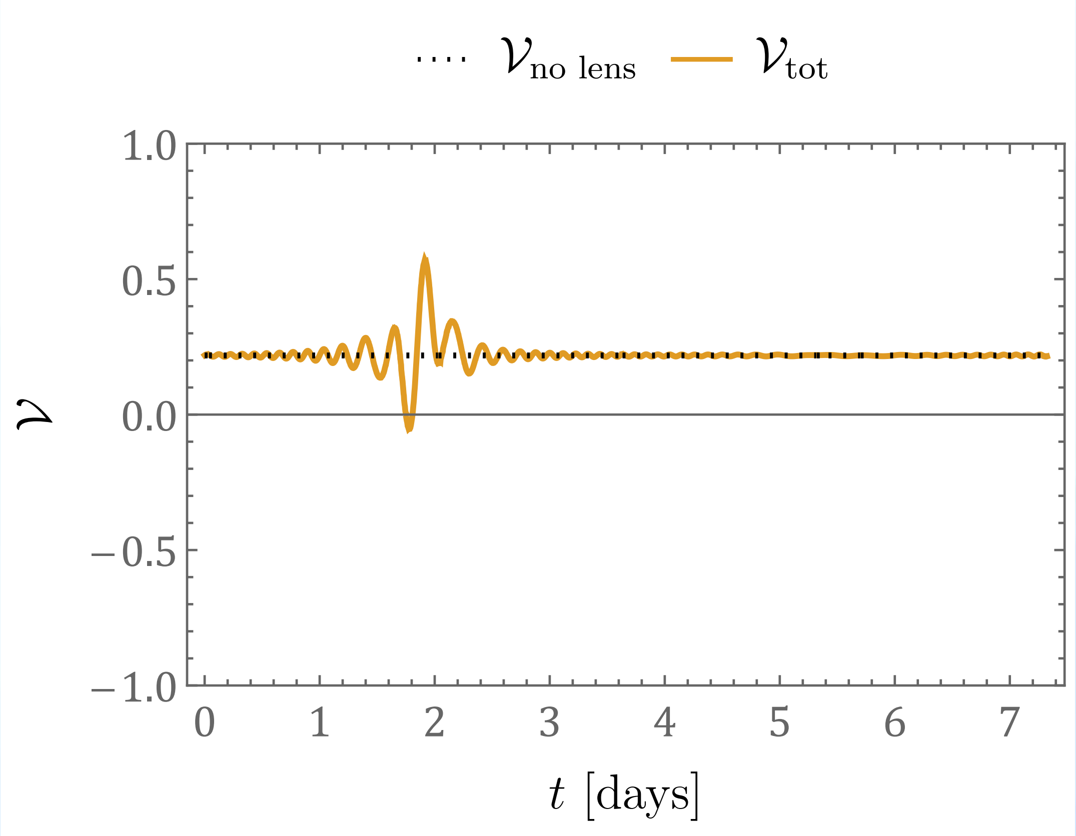

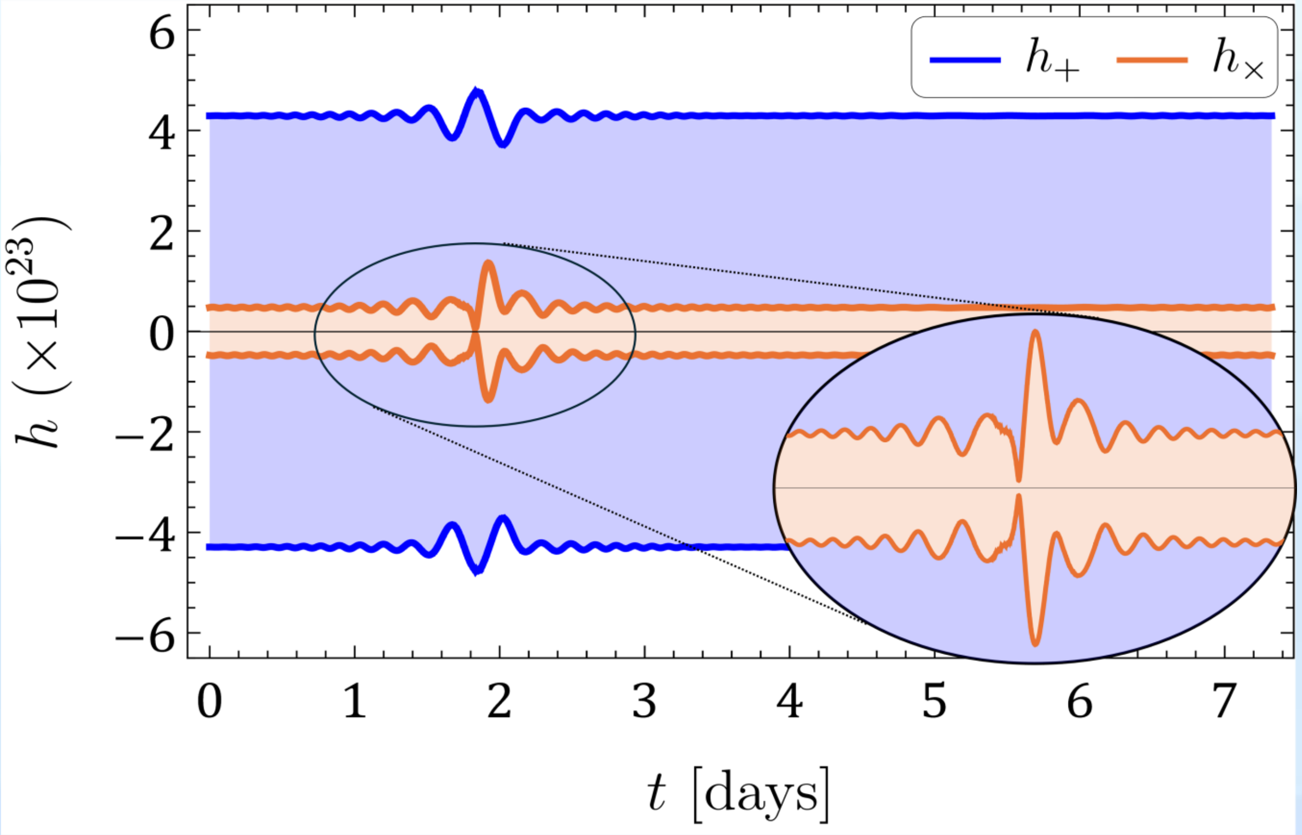

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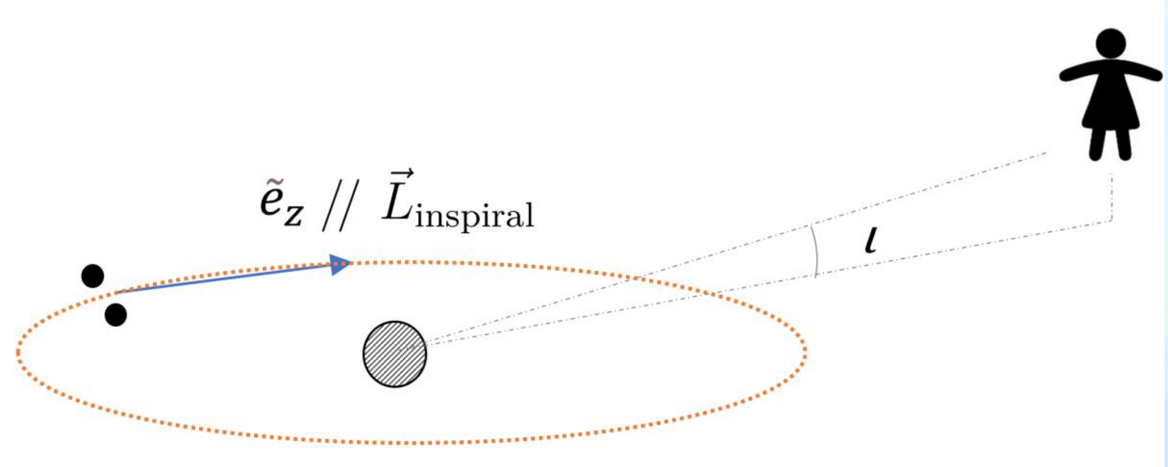
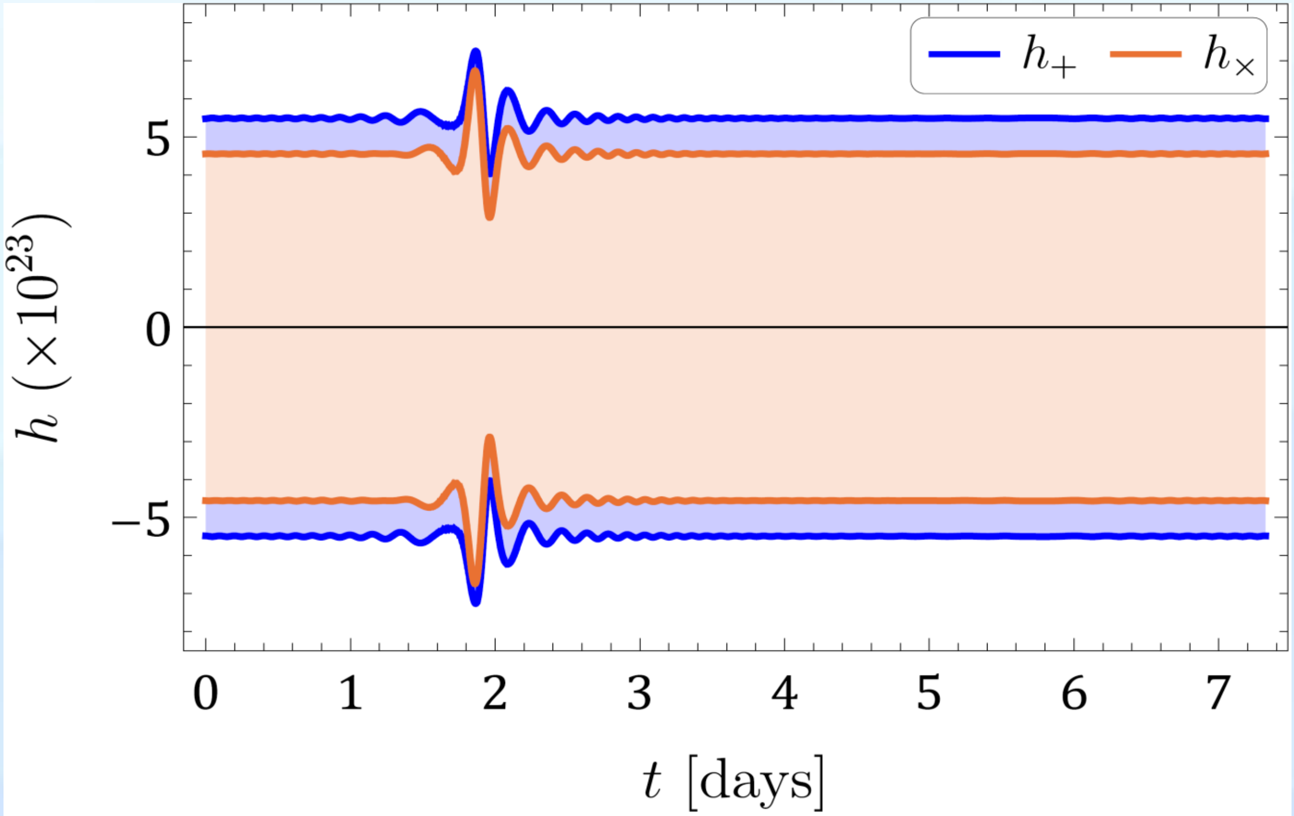


(not to scale)



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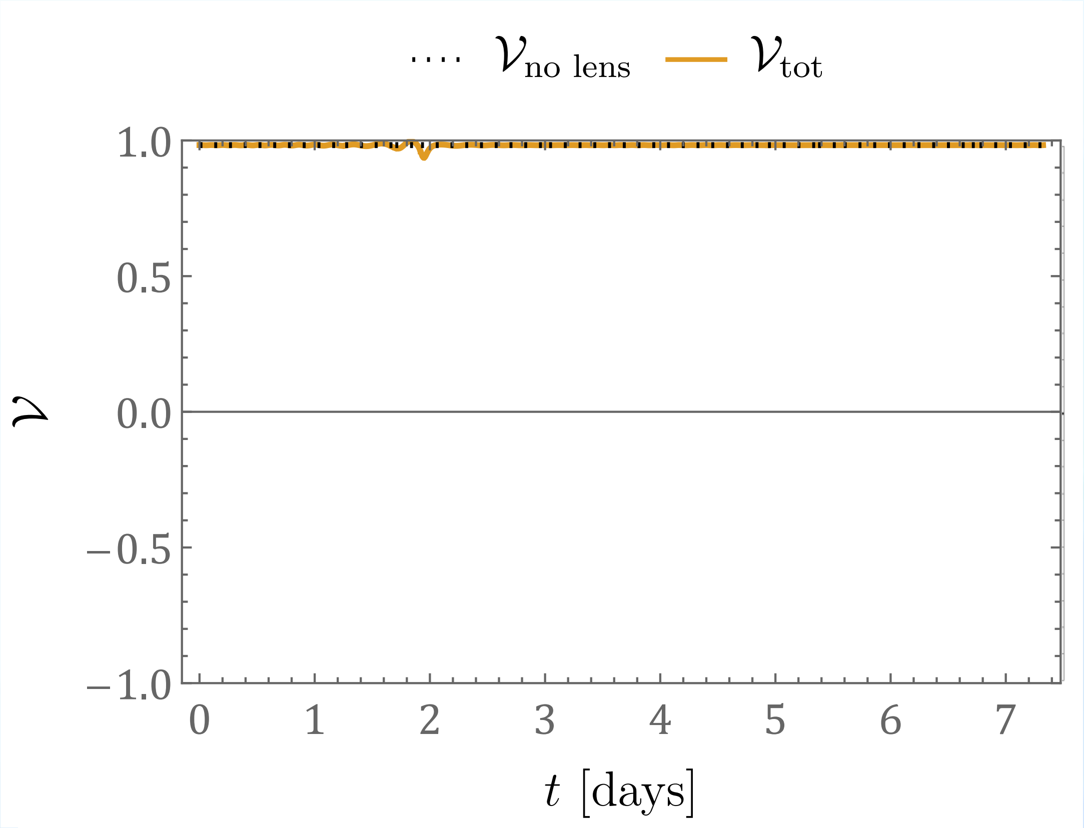
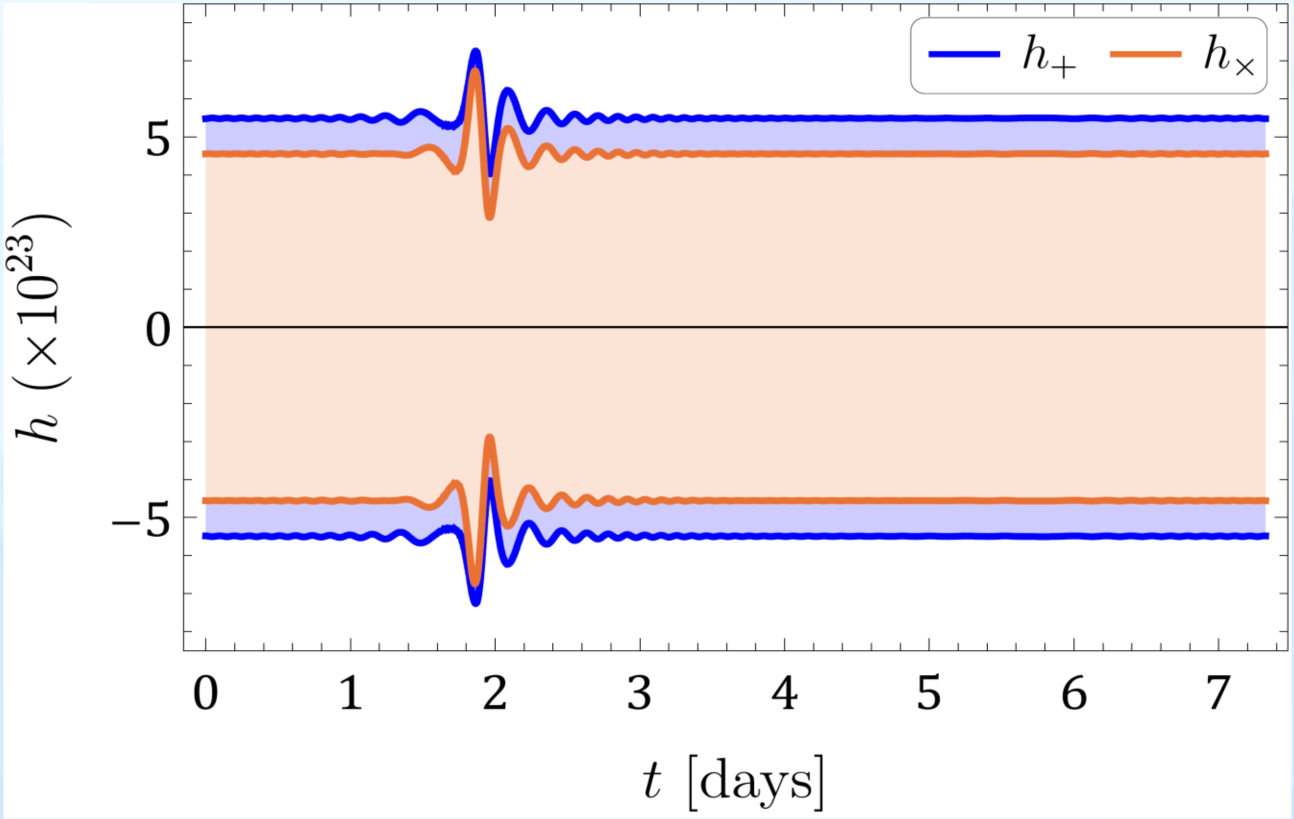
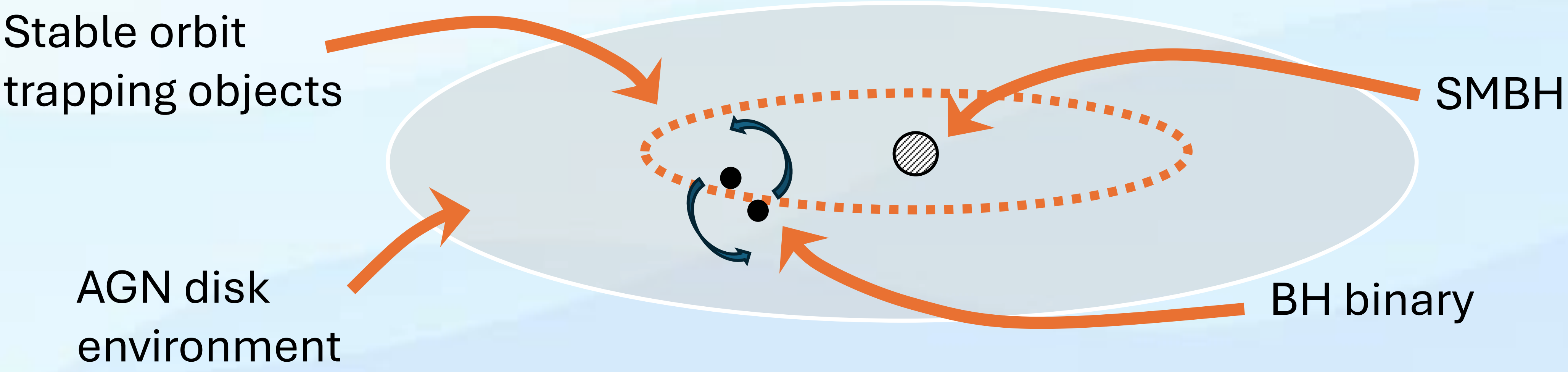




Hierarchical triple systems : GW190521 ?

- 34 days later: AGN signal from the “same” location
- Suspect merger ocurred $\sim 350\ r_S$ away from AGN : triple system



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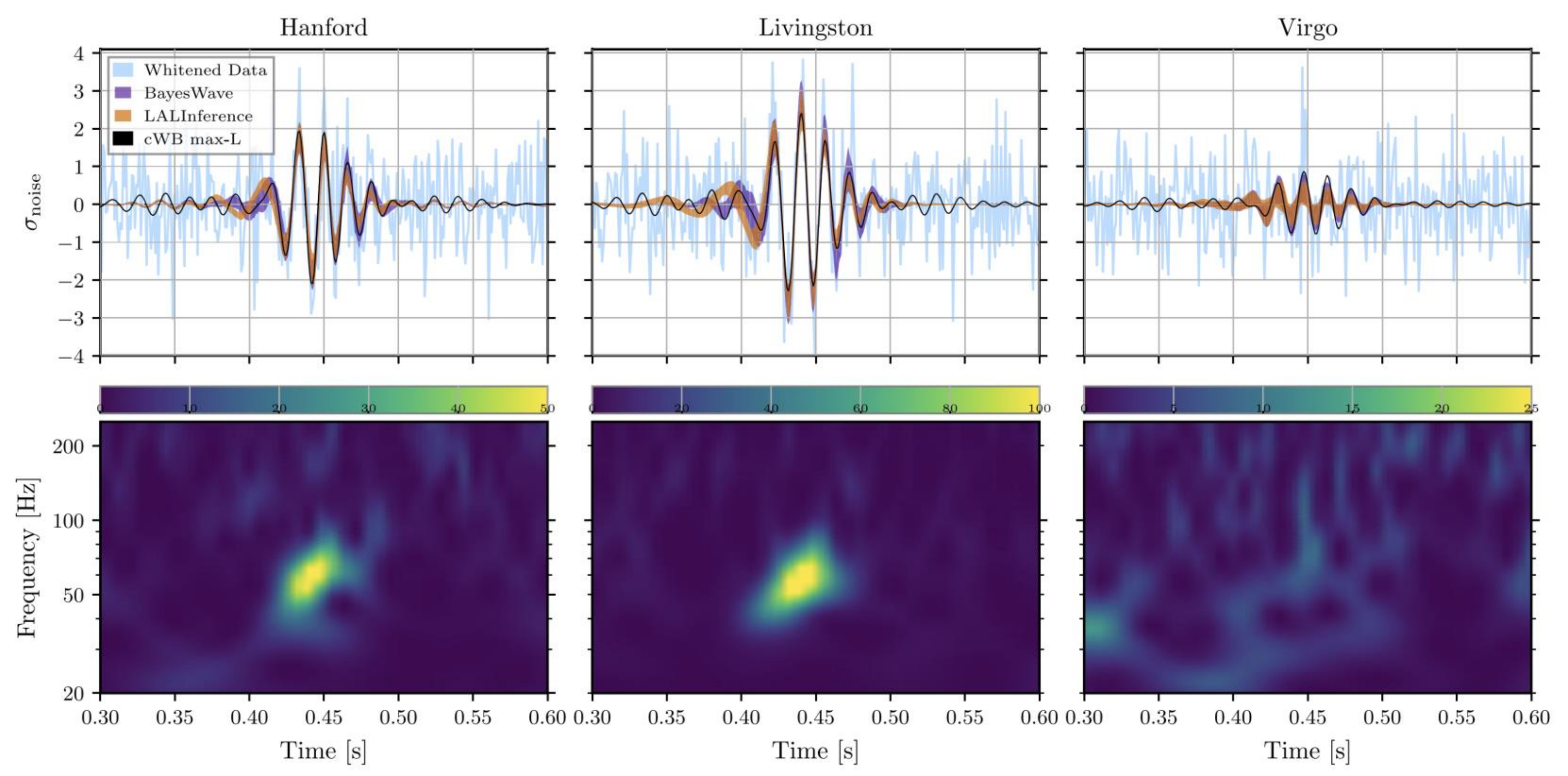
Considering **both GW & EM** signals, **evidence** for hierarchical triple:

marginal Ashton *et al.* (2021), Palmese *et al.* (2021)

confident Graham *et al.* (2020), Morton *et al.* (2023)

Hierarchical triple systems : GW190521 ?

Observational fact: 5 years ago ...



Best fit masses:

$$m_1 = 85M_\odot,$$
$$m_2 = 66M_\odot$$

(heavy !)

R. Abbott *et al.* (LIGO Scientific Collaboration and Virgo Collaboration),
GW190521: A Binary Black Hole Merger with a Total Mass of 150 $M_\odot$
Phys. Rev. Lett. **125**, 101102, 2020

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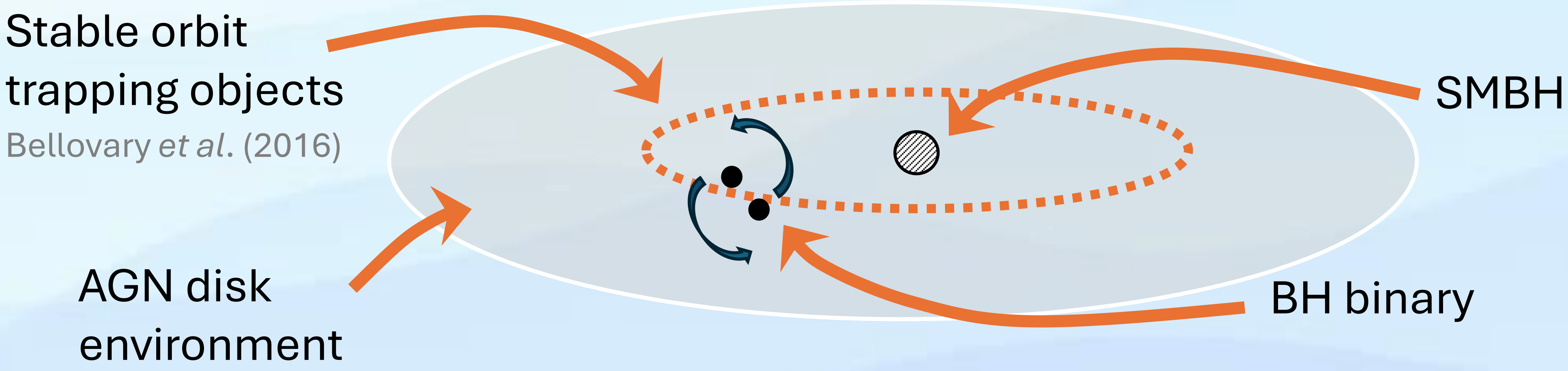

GW lensing: wave optics

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At this frequency, **low mass AGN** with $M \sim \mathcal{O}(10^6) M_\odot$ fulfil the wave optics requirement

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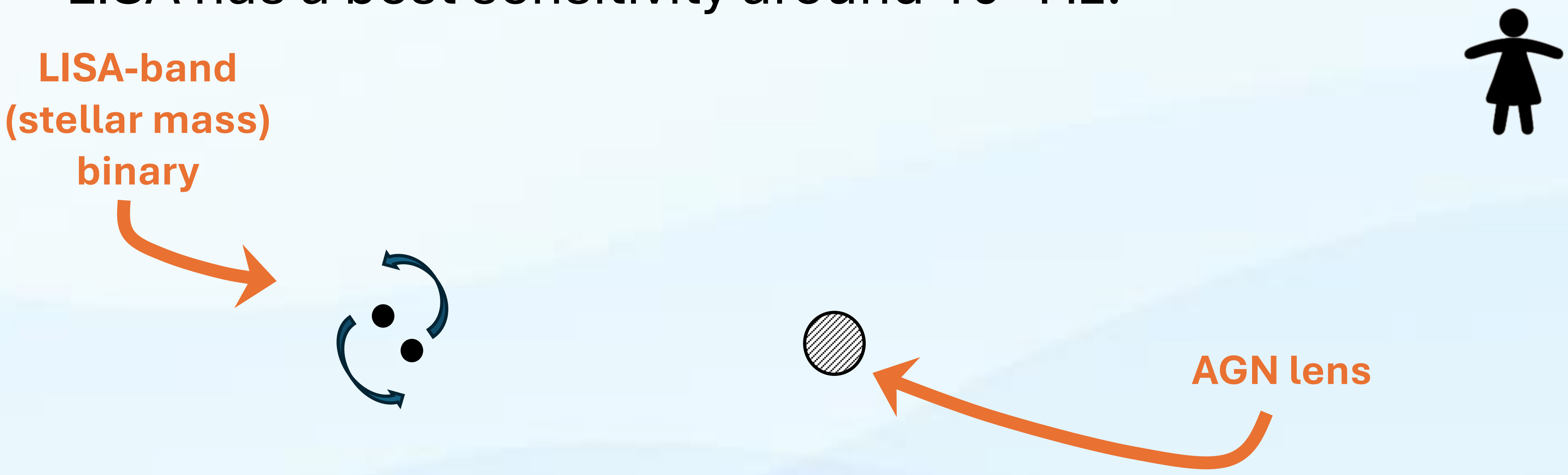

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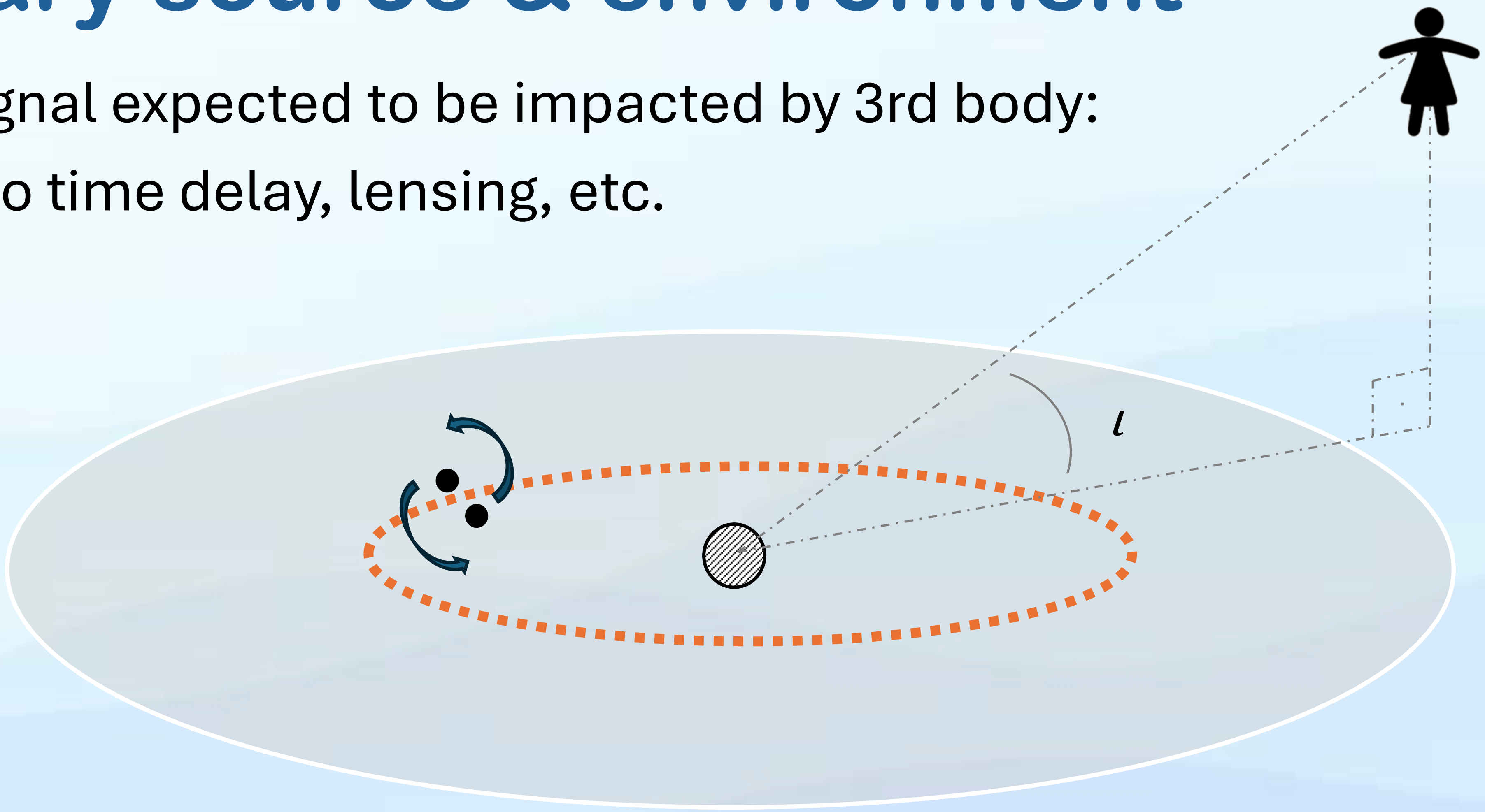
How to capture wave effects ?

Hierarchical triple systems : binary source & environment

GW signal expected to be impacted by 3rd body:

Shapiro time delay, lensing, etc.



GW lensing: wave optics

BH lenses, historical works, at the formal level :

- Matzner (1968)
- Peters (1976)
- Chrzanowski *et al.* (1976)
- De Logi, Kovacs (1977)
- Futterman *et al.* (1988)

…

More recently: Dolan (2018)

GW lensing: wave optics

Reference work for phenomenology :

Wave effects in gravitational lensing of gravitational waves from chirping binaries

Ryuichi Takahashi (Kyoto U.), Takashi Nakamura (Kyoto U.)

May, 2003

28 pages

Published in: *Astrophys.J.* 595 (2003) 1039-1051

e-Print: astro-ph/0305055 [astro-ph]

DOI: 10.1086/377430

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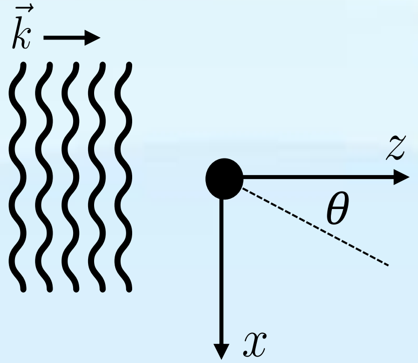
Polarisation

Quantifying the signal polarisation content $\mathcal{V} \in [-1, 1]$:

$$\mathcal{V} \equiv \frac{2\mathrm{Im}[\tilde{h}_+ \tilde{h}_\times^*]}{|\tilde{h}_+|^2 + |\tilde{h}_\times|^2} = V/I$$



Wave optics lensing is **polarisation dependent**, e.g. :

$$\frac{\mathrm{d}\sigma}{\mathrm{d}\Omega} = M^2 \frac{\cos^8\left(\frac{\theta}{2}\right) + \sin^8\left(\frac{\theta}{2}\right)}{\sin^4\left(\frac{\theta}{2}\right)} + 2M^2\sqrt{1 - \mathcal{V}_{\text{incident}}^2}\,\cos^4\left(\frac{\theta}{2}\right)\cos(4\phi)$$

Tensorial wave optics : BHPT

Technicality : in principle, should sum $\lim_{r\to\infty} \sum_{\ell m} \Psi_{\bullet}^{\ell m}$

In practice : $\sum_{\ell m} \lim_{r\to\infty} \Psi_{\bullet}^{\ell m}$

Derive final

$h_{\mu\nu}$

(incl. lensing)

$\longleftarrow$

Solve differential equ. such that

$\Psi_{\bullet}^{\ell m} \sim \Psi_{\bullet}^{\text{source}} + \Psi_{\bullet}^{\text{lensed}}$

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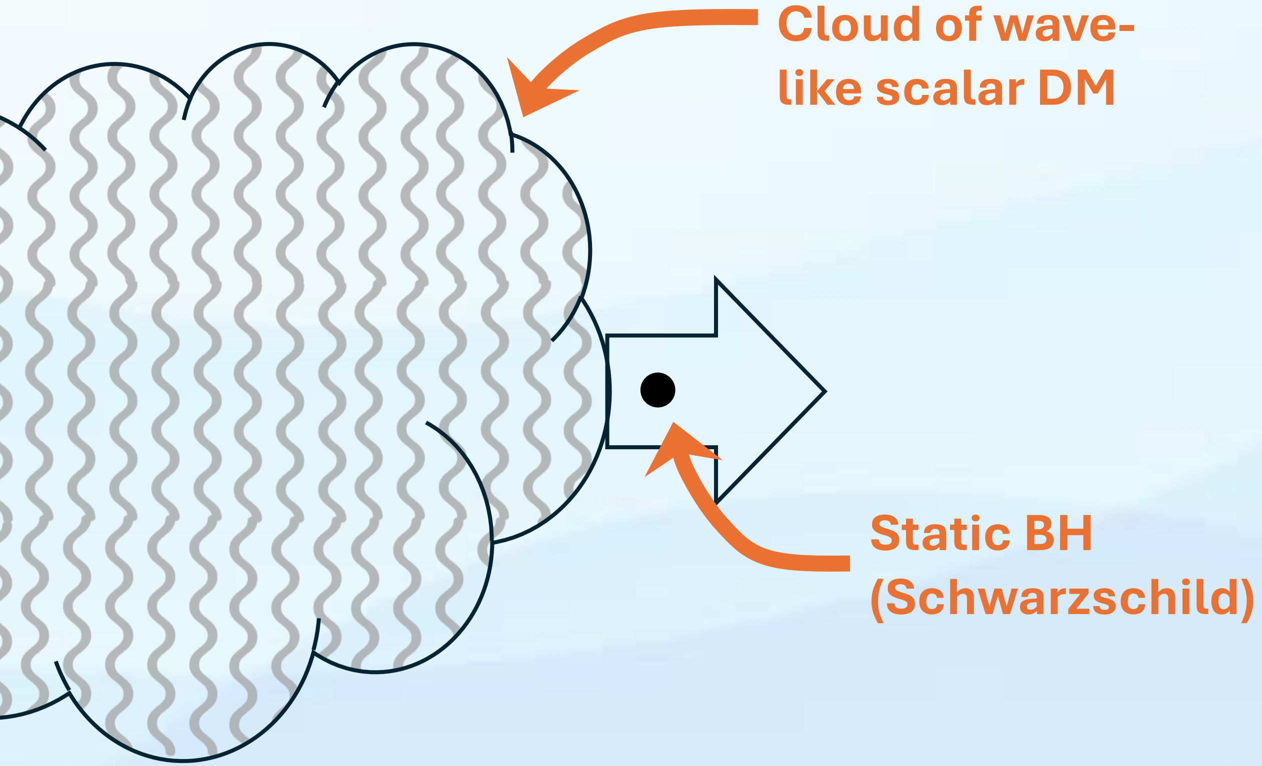
In practice : $\sum_{\ell m} \lim_{r\to\infty} \Psi_\bullet^{\ell m}$ ... **diverges** analytically & numerically

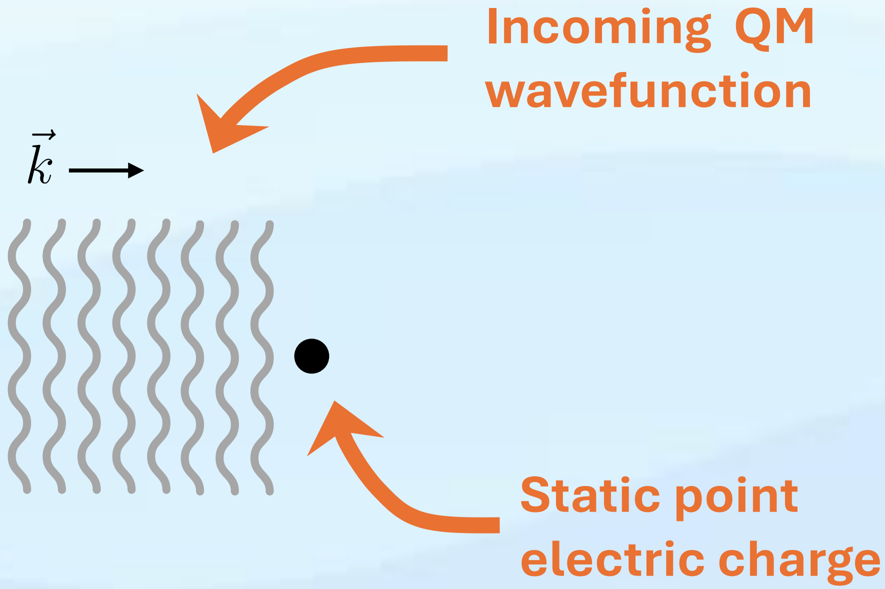
Derive final $h_{\mu\nu}$ (incl. lensing) $\longleftarrow$ **non standard summation methods** — Solve differential equ. such that $\Psi_\bullet^{\ell m} \sim \Psi_\bullet^{\text{source}} + \Psi_\bullet^{\text{lensed}}$

Insights from quantum mechanics

GR scalar wave, NR simulation



QM, analytical solution



Insights from quantum mechanics

If one doesn't know the exact solution :

Possible to solve the differential equation in multipole space, by

- Taking $kr \gg 1$ limit
- Requiring $\Psi \sim \Psi_{\text{plane}} + \Psi_{\text{spherical}}$

$\rightarrow$ **Correctly** recover $\Psi_{\text{spherical}}$ = $-\dfrac{\tilde{\gamma}}{kr(1-\cos(\theta))}\, e^{ikr - i\gamma \ln kr(1-\cos(\theta))}$

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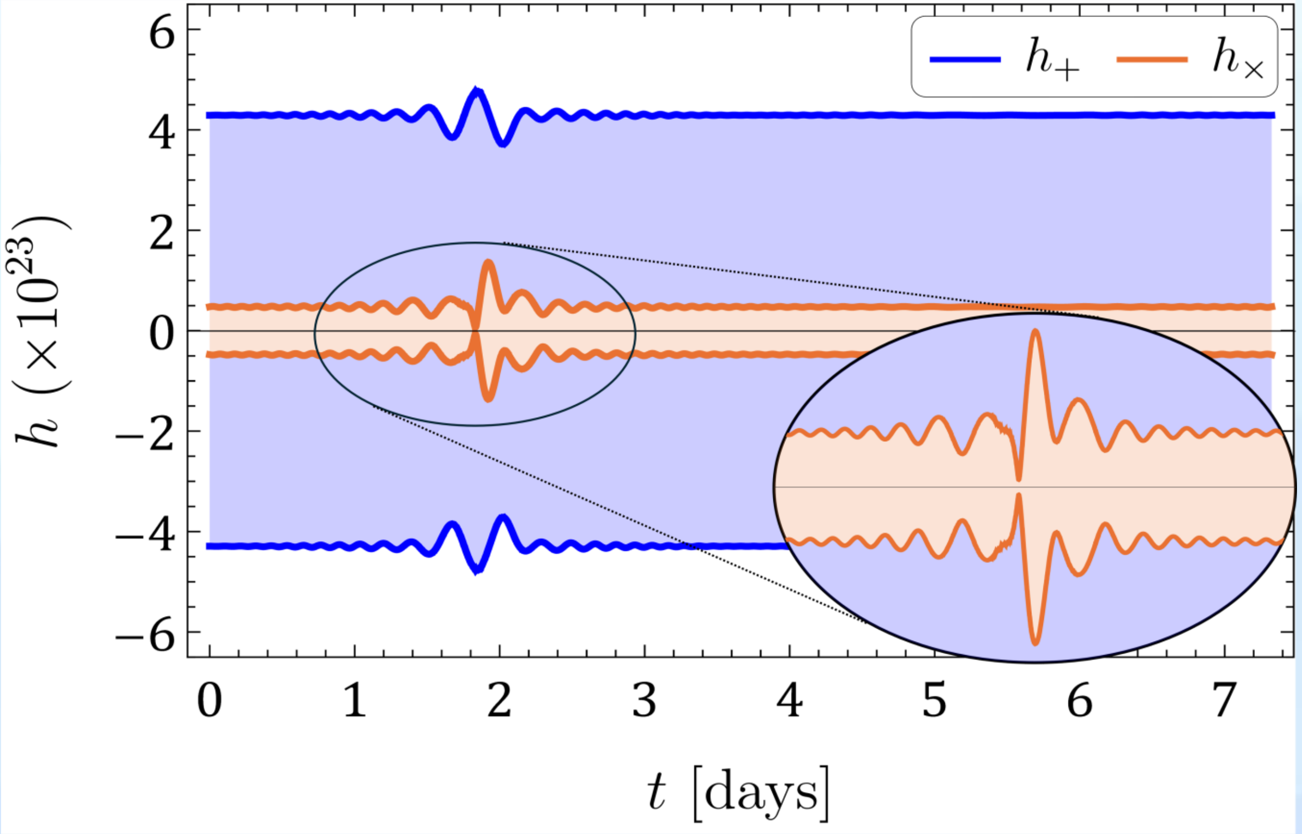
Wave optics lensing in triple systems: towards a phenomenology

Toy GW190521-inspired source, in LISA- « optimal » wave optics

Total waves envelopes, for both polarisations



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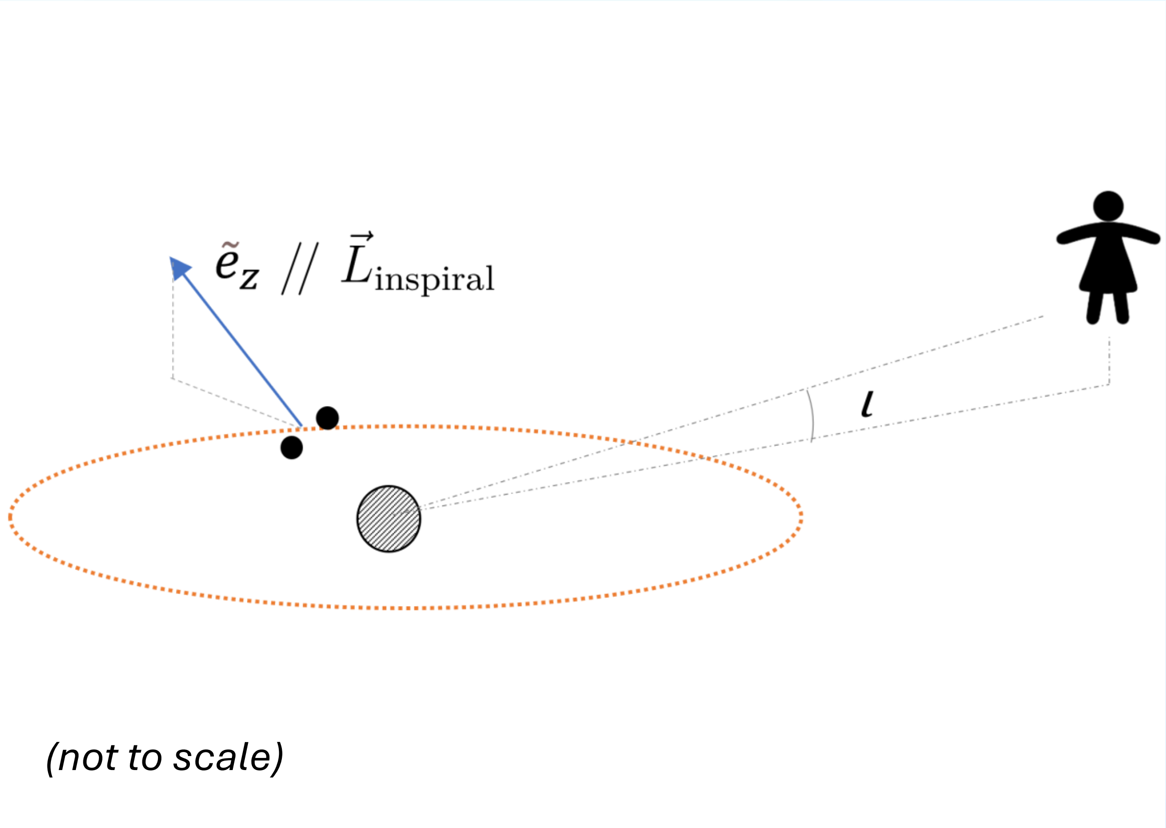

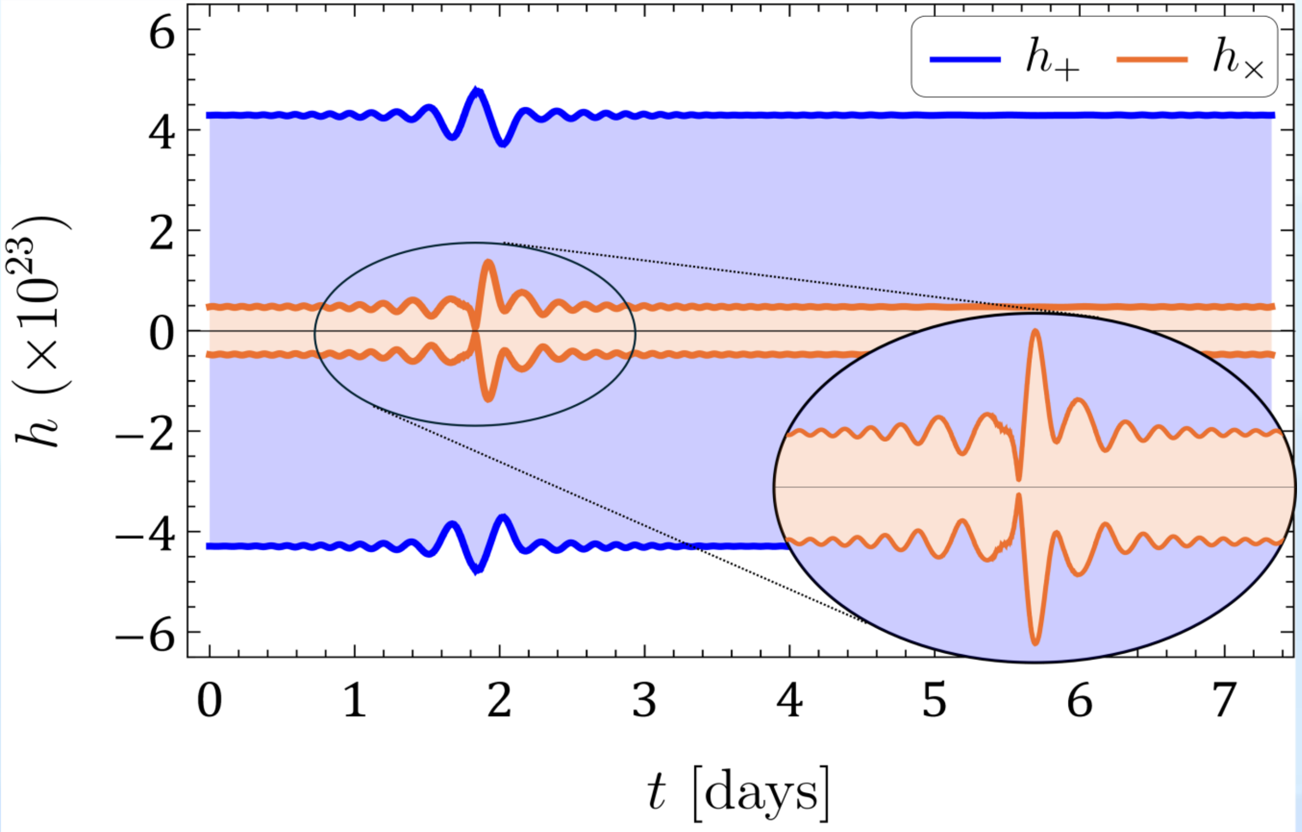

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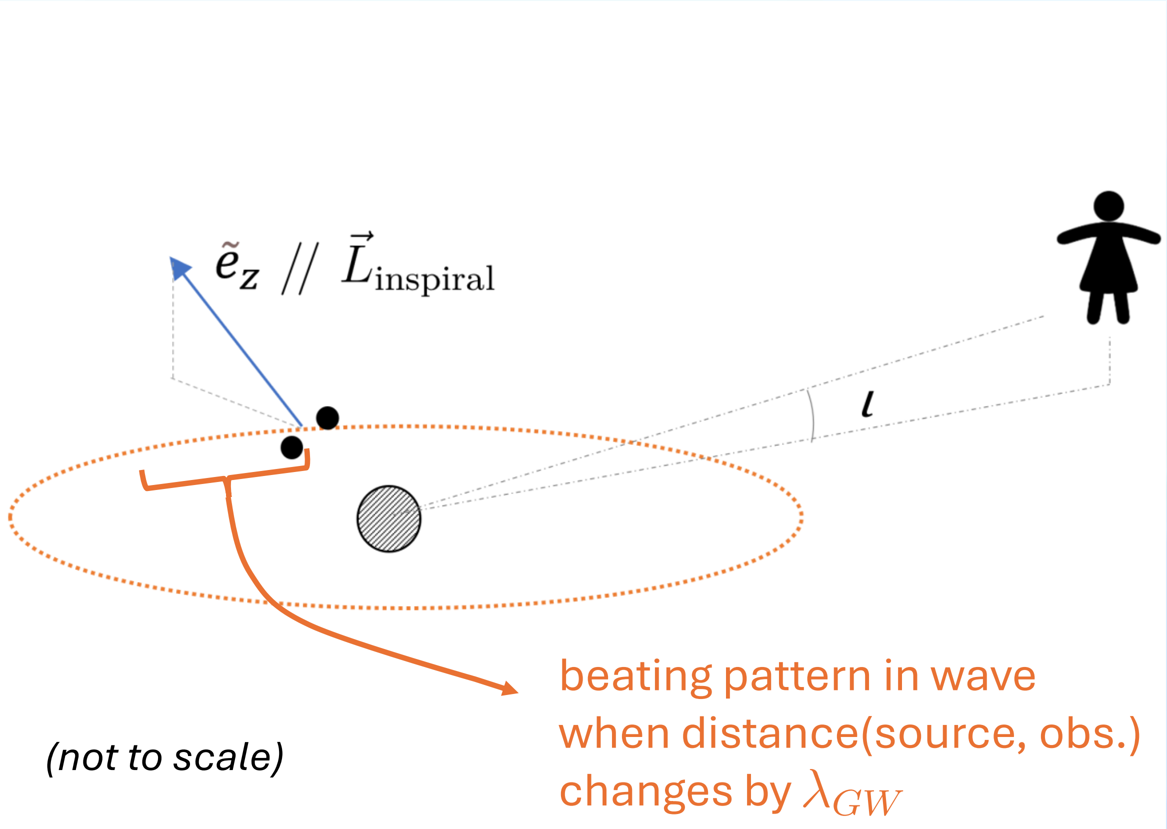

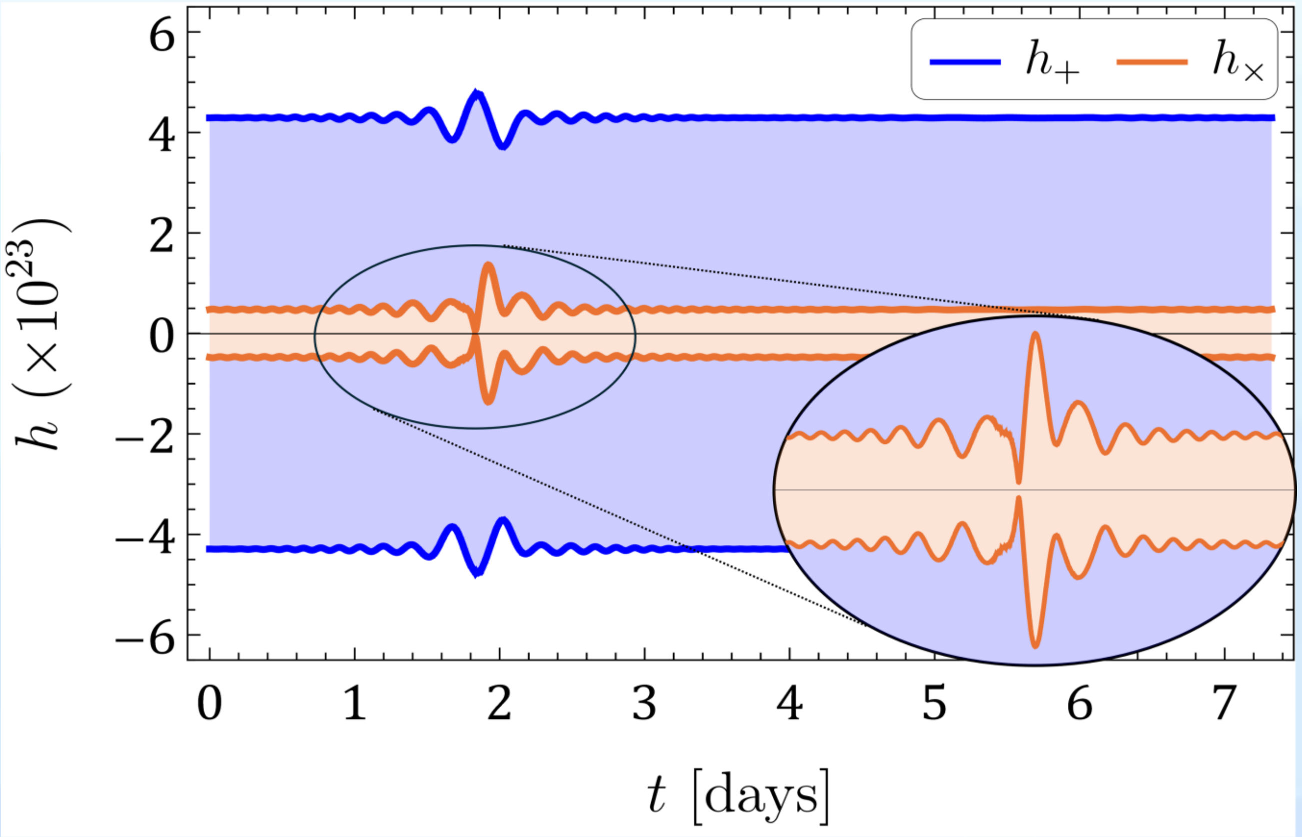

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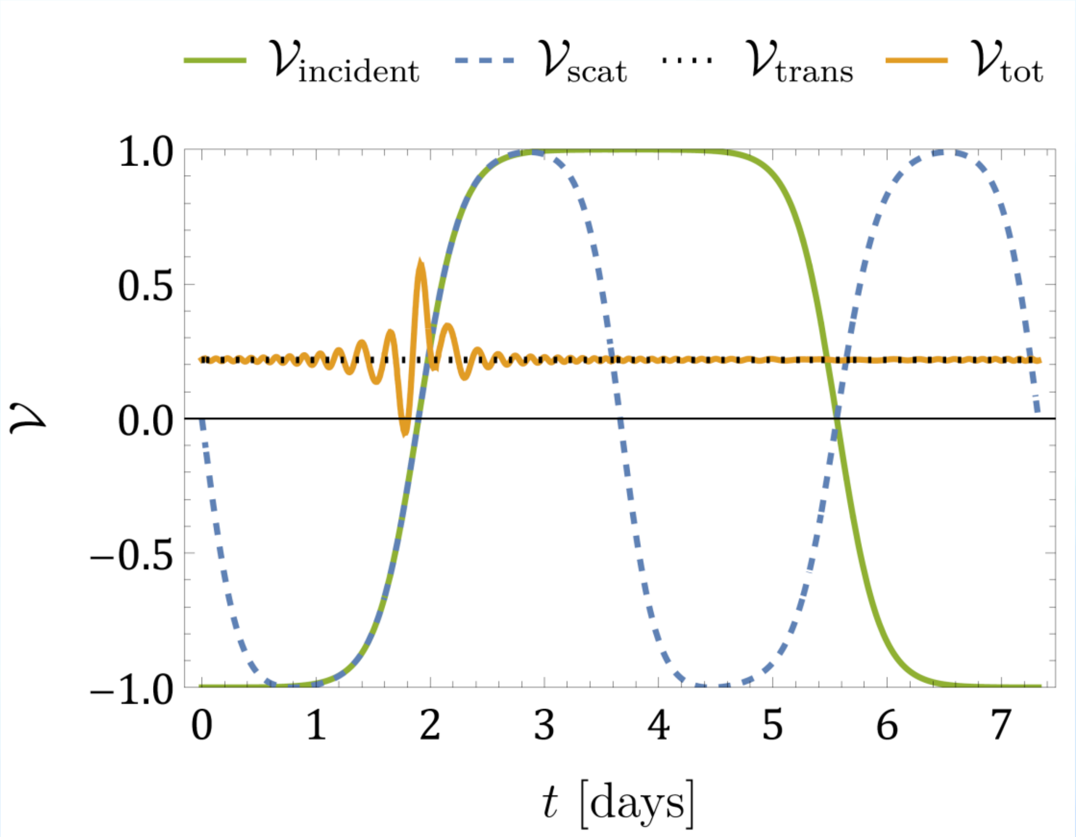

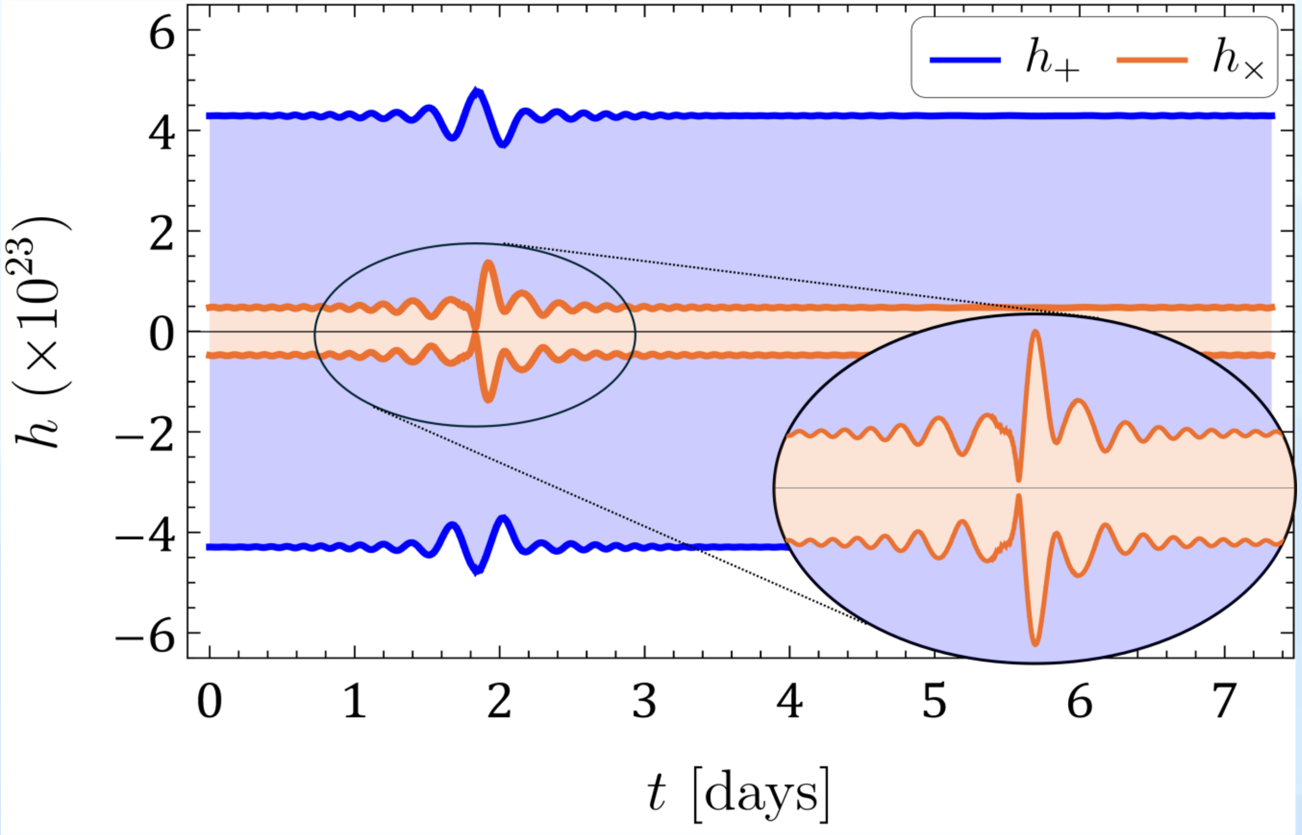

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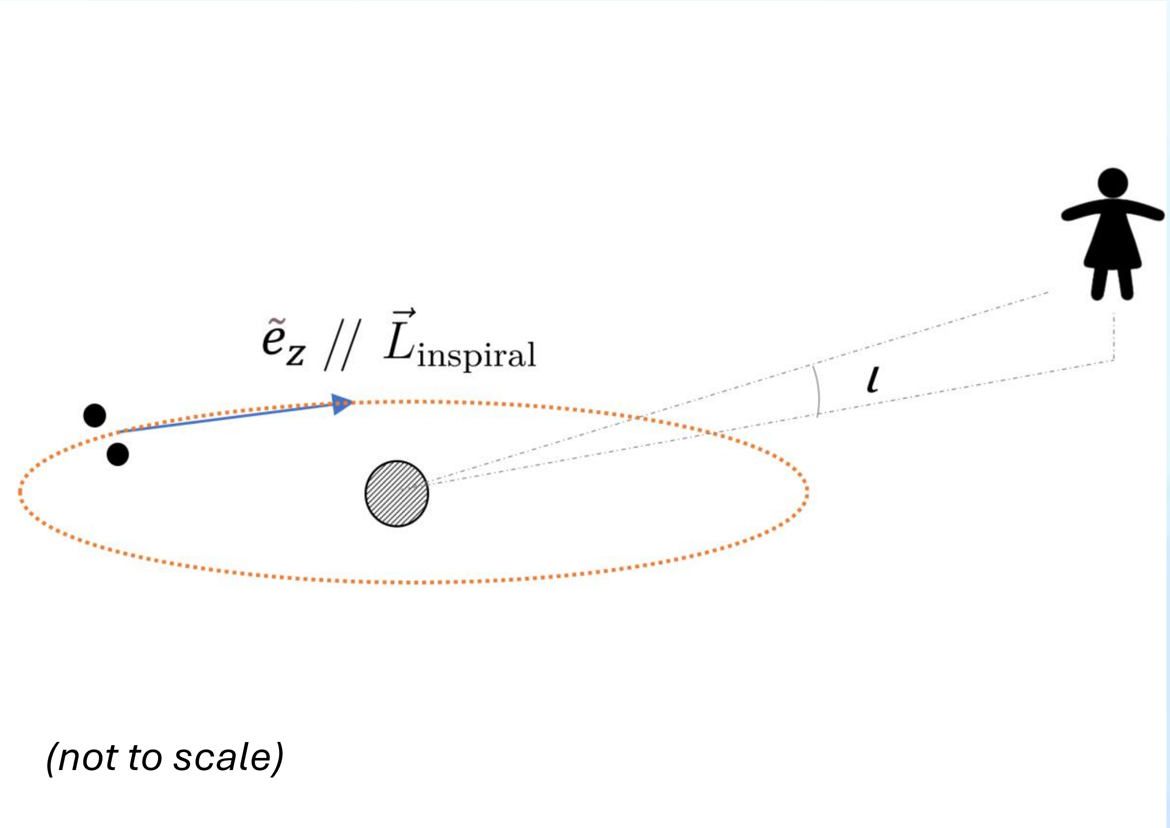

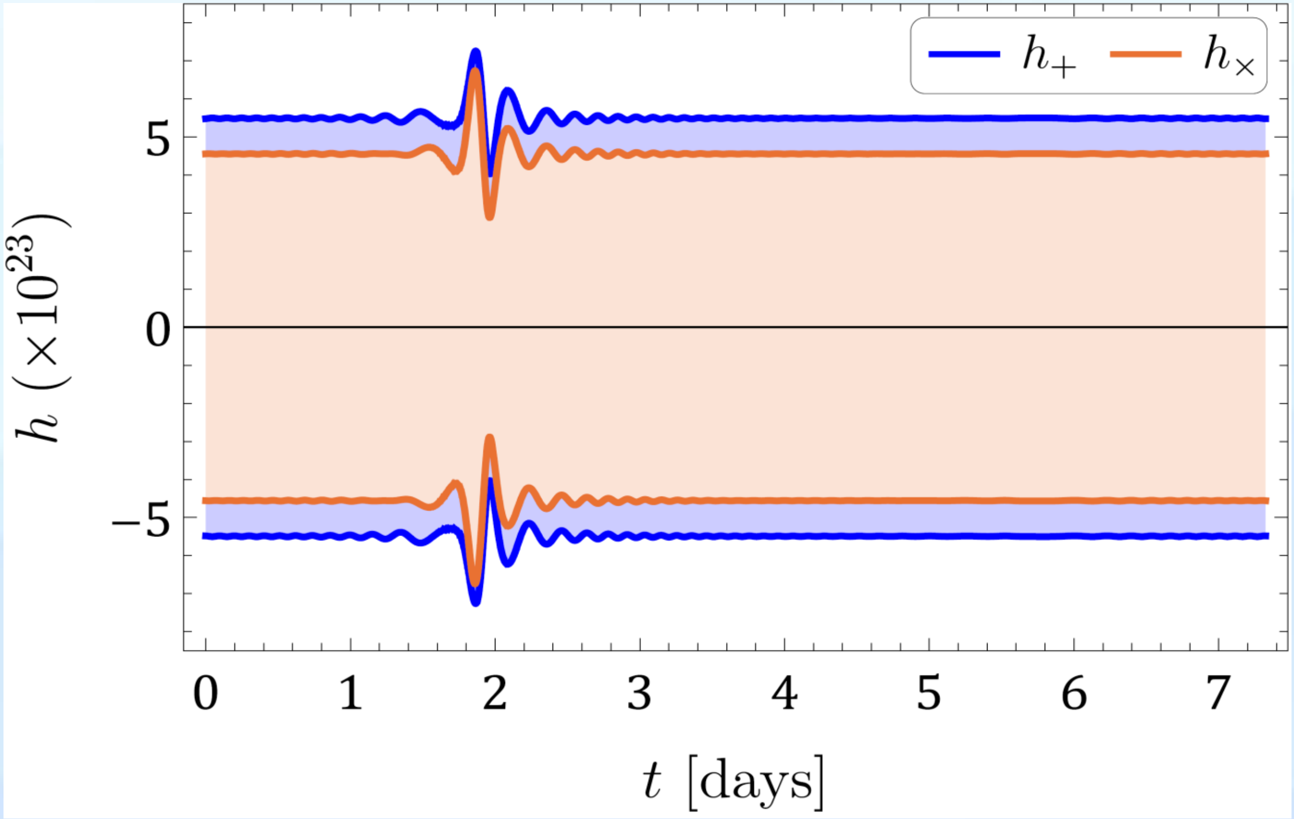

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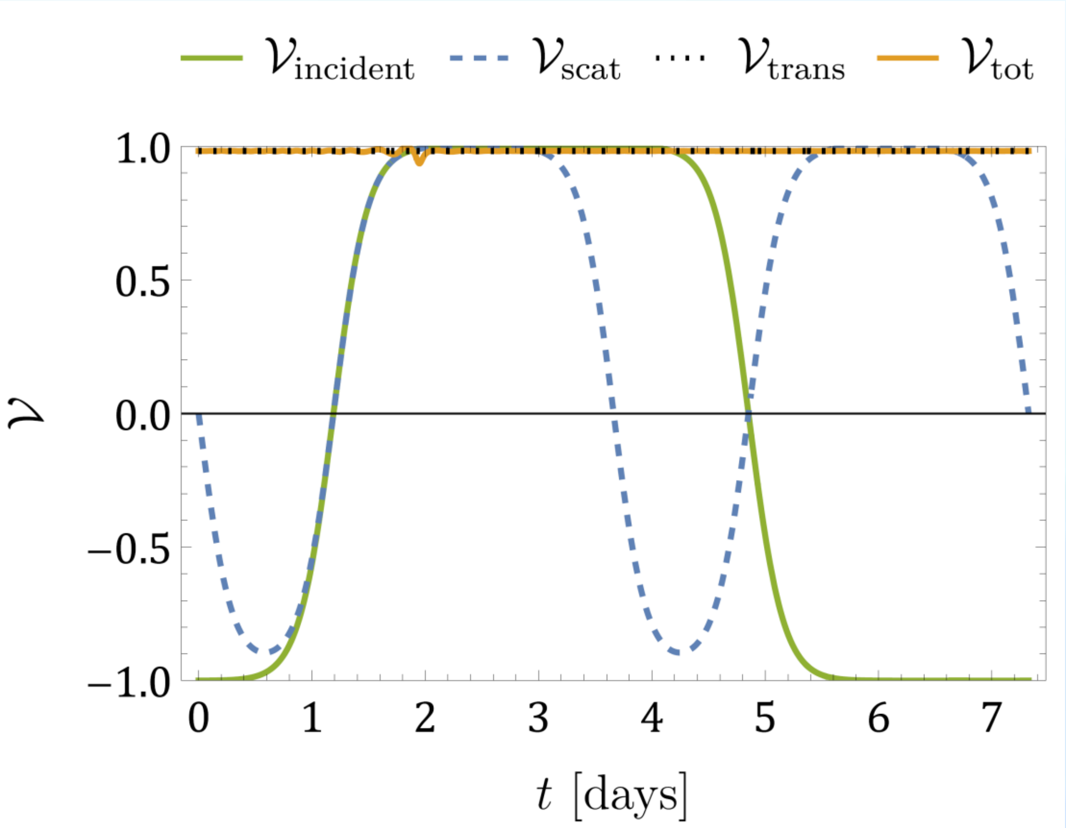
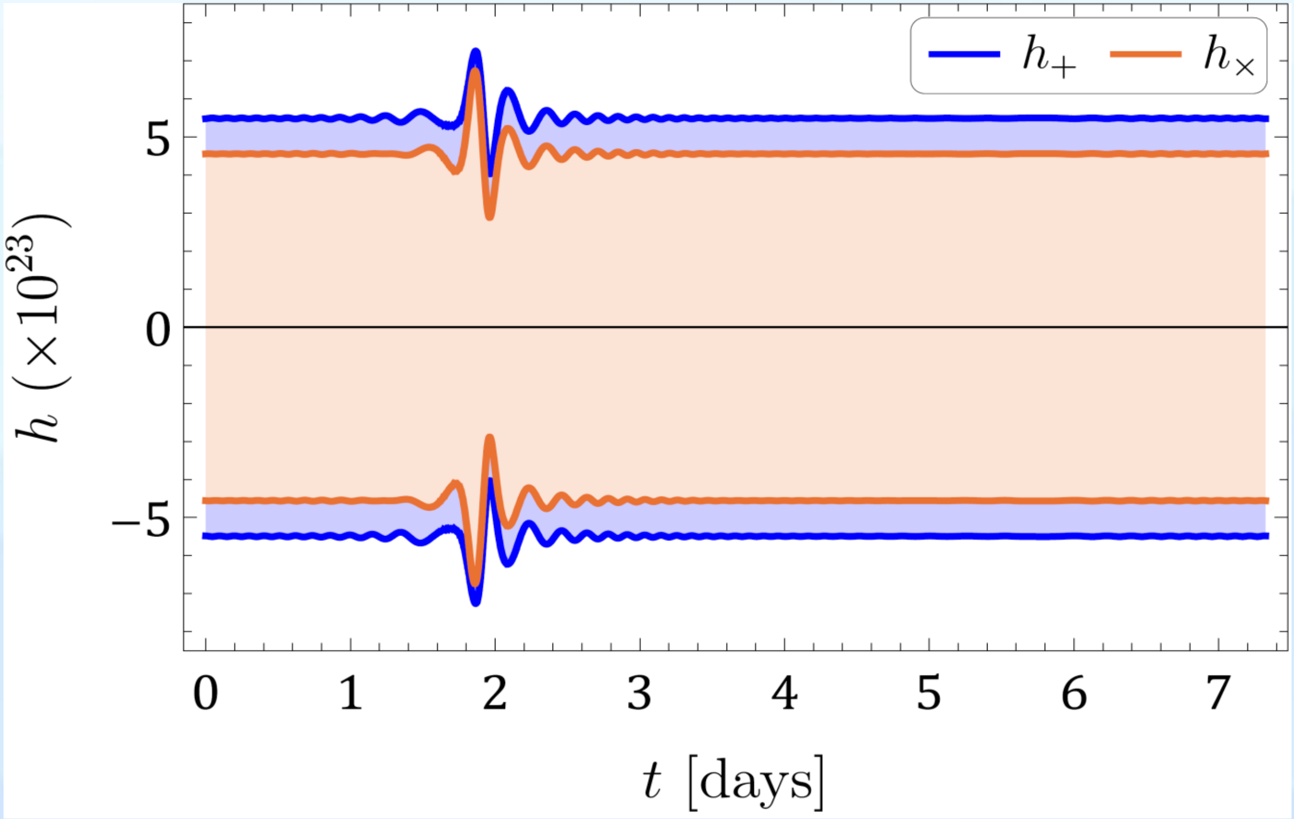
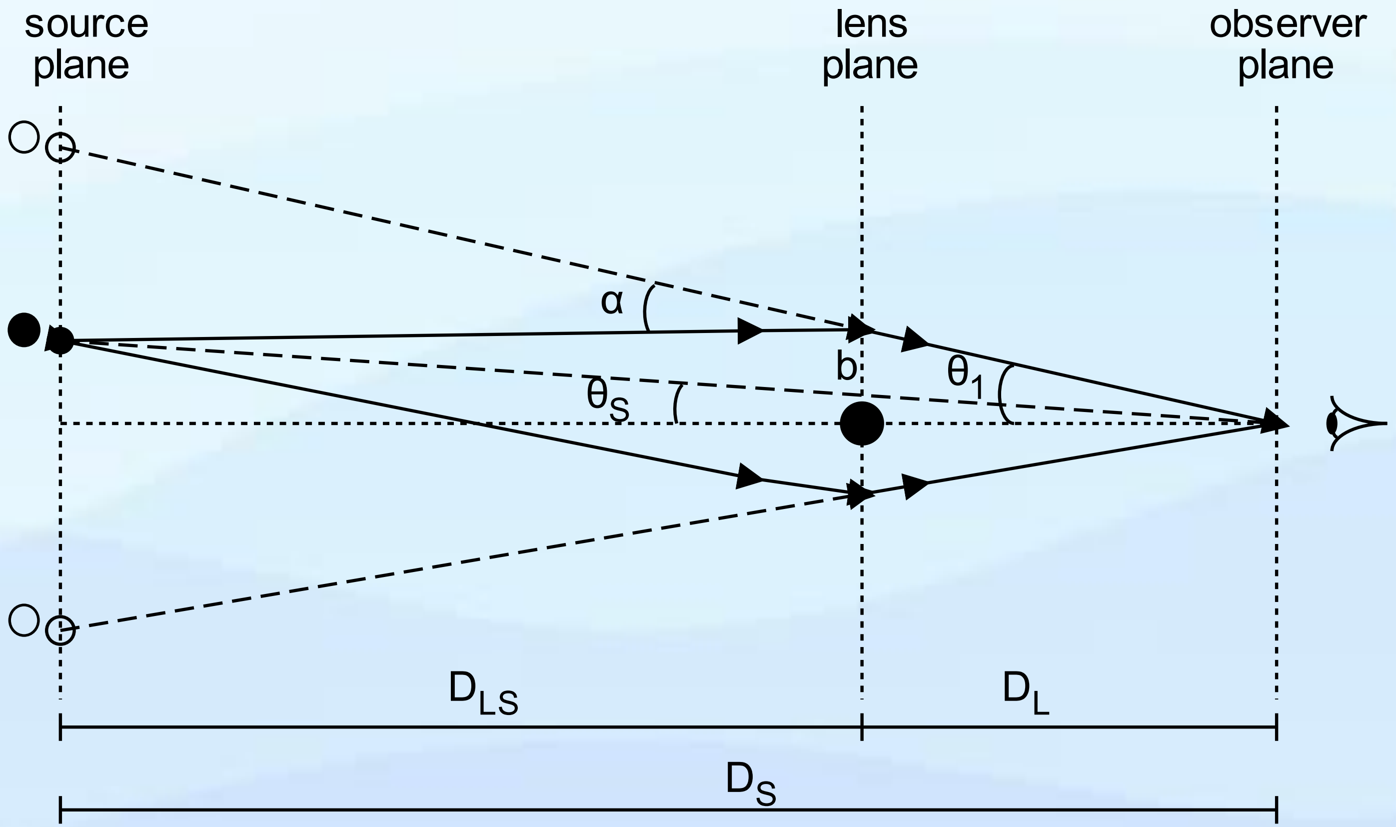




GW lensing: geometrical optics

Assume you have a notion of ***ray*** :

→ Usual lensing picture (deflection angle, etc.)



GW lensing: geometrical optics

Conceptually similar to (my) undergraduate lab optics