MAPPING OUT THE DARK MATTER IN THE MILKY WAY

Lina Necib

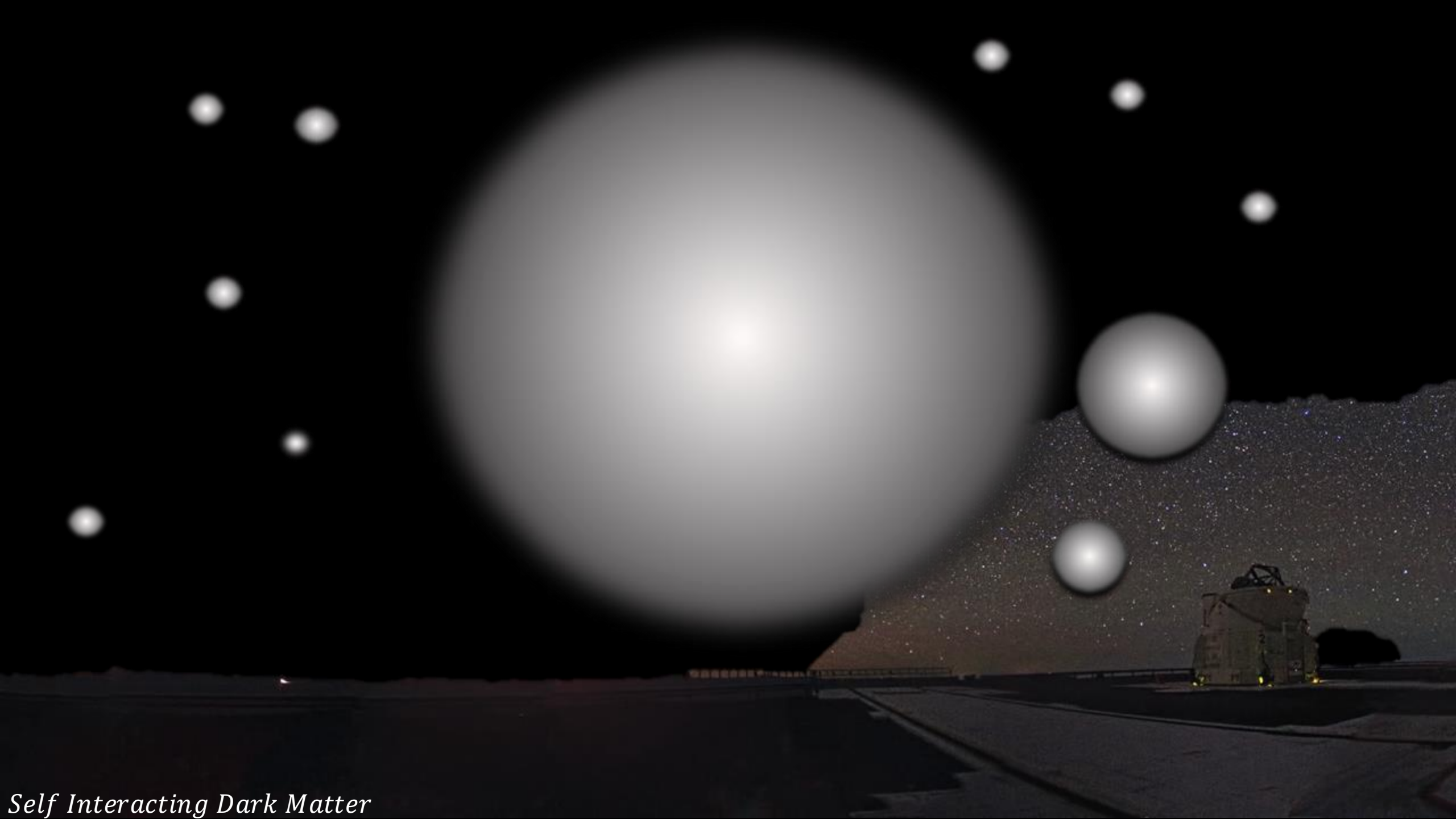
Massachusetts Institute of Technology

MIT KAVLI INSTITUTE

Credit: ESO/Y. Beletsky


ΛCDM


Warm Dark Matter


Self Interacting Dark Matter

Something else altogether


ΛCDM

GOAL:
MAPPING FROM THE STARS TO THE DARK MATTER

GOAL:
MAPPING FROM THE STARS TO THE DARK MATTER

COSMIC CODEBREAKERS

Streams

Nora Shipp

After MIT: Faculty at University of Washington

Nathaniel Starkman

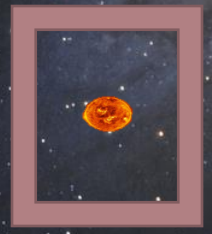
Brinson Fellow

Galactic Center/Halo

Xiaowei Ou

Abdelaziz Hussein

Solar Neighborhood



Zeineb Mezghanni

Xiuyuan Zhang

Dwarf Galaxies

Tri Nguyen, PhD' 24

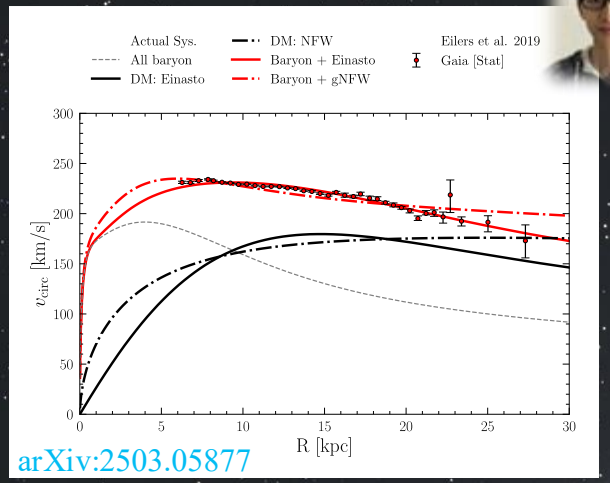
After MIT: CIERA Fellow, Northwestern

Divij Sharma

Lina Necib, MIT

TODAY'S TALK: MAPPING THE DARK MATTER DENSITY PROFILE IN THE MILKY WAY AND DWARF GALAXIES

Xiaowei Ou

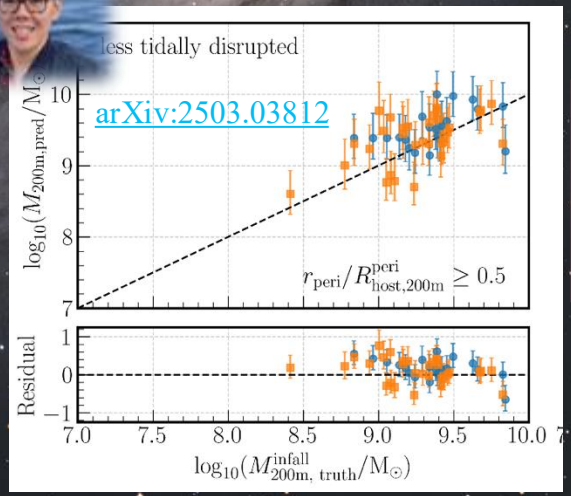


arXiv:2503.05877

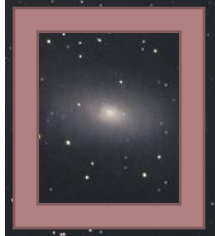
Used synthetic surveys to build the true systematic uncertainties of the circular velocity measurements of the Milky Way, and obtained a new Dark Matter density profile.

Galactic Center/Halo

Tri Nguyen

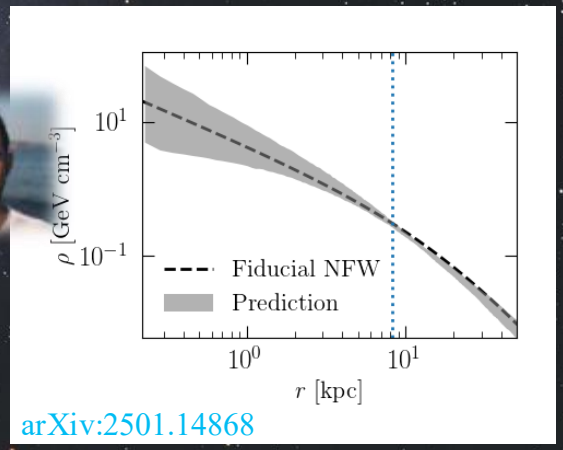


Dwarf Galaxies



Used Graph Neural Networks to construct the density profiles of Dark Matter in dwarf galaxies, based on line-of-sight velocities.
Tested this new method on FIRE simulations to show robustness. Next we will apply this to data.

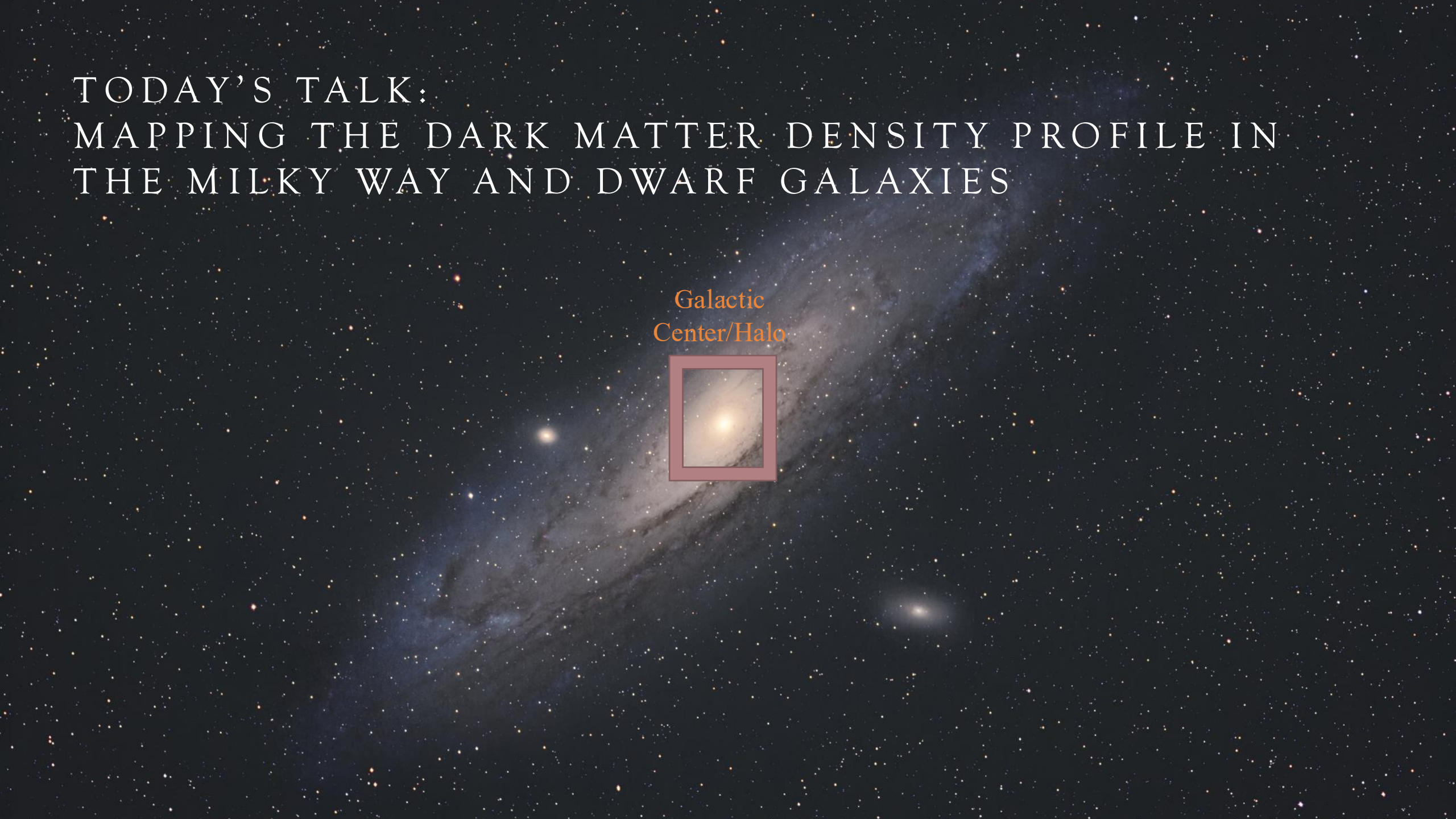
Abdelaziz Hussein



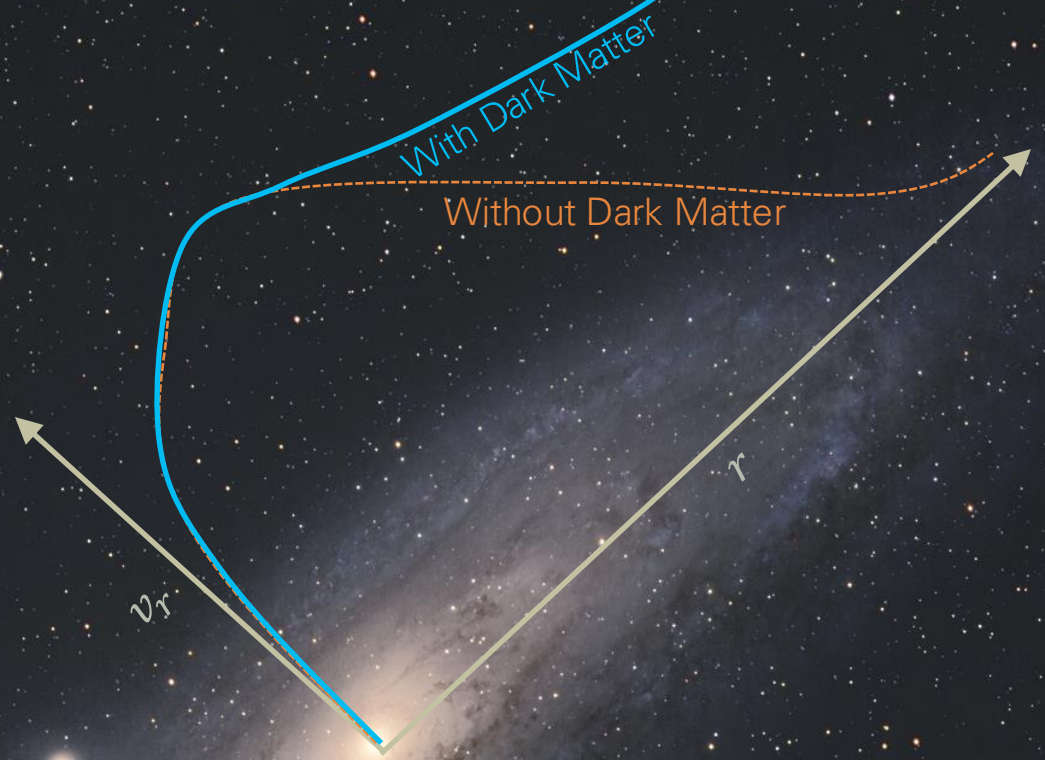
arXiv:2501.14868

Built a theoretical range for the density profile of Dark Matter at the center of the Milky Way, based on adiabatic contraction and baryonic feedback.

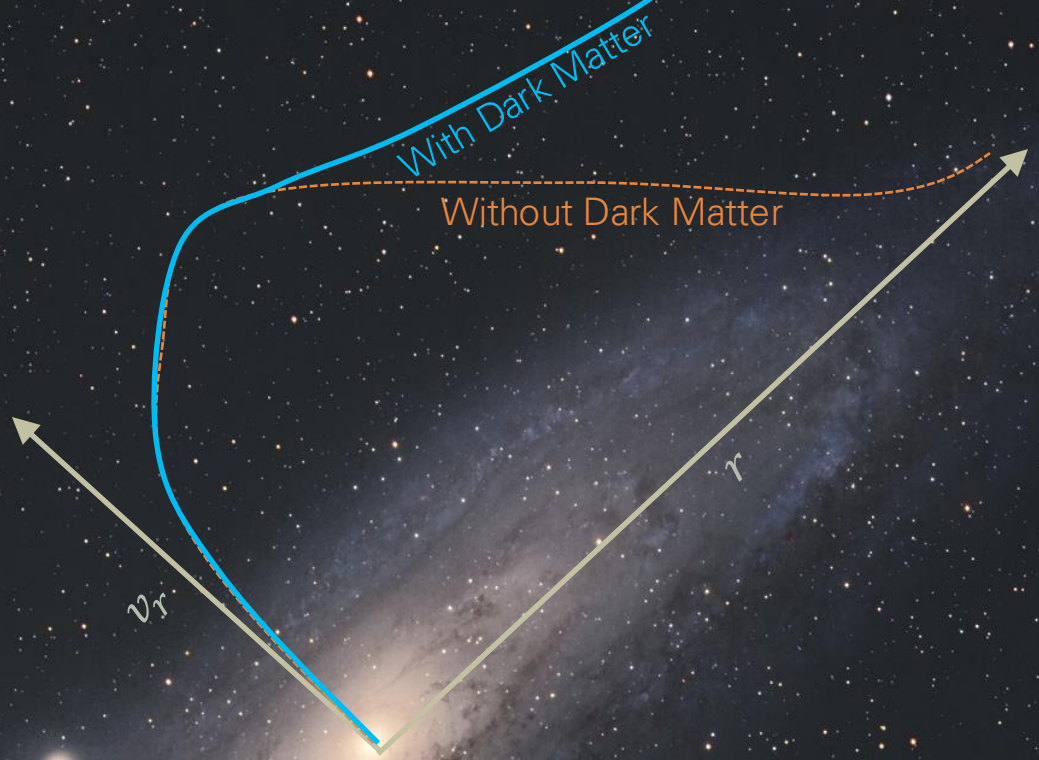
TODAY'S TALK:
MAPPING THE DARK MATTER DENSITY PROFILE IN THE MILKY WAY AND DWARF GALAXIES



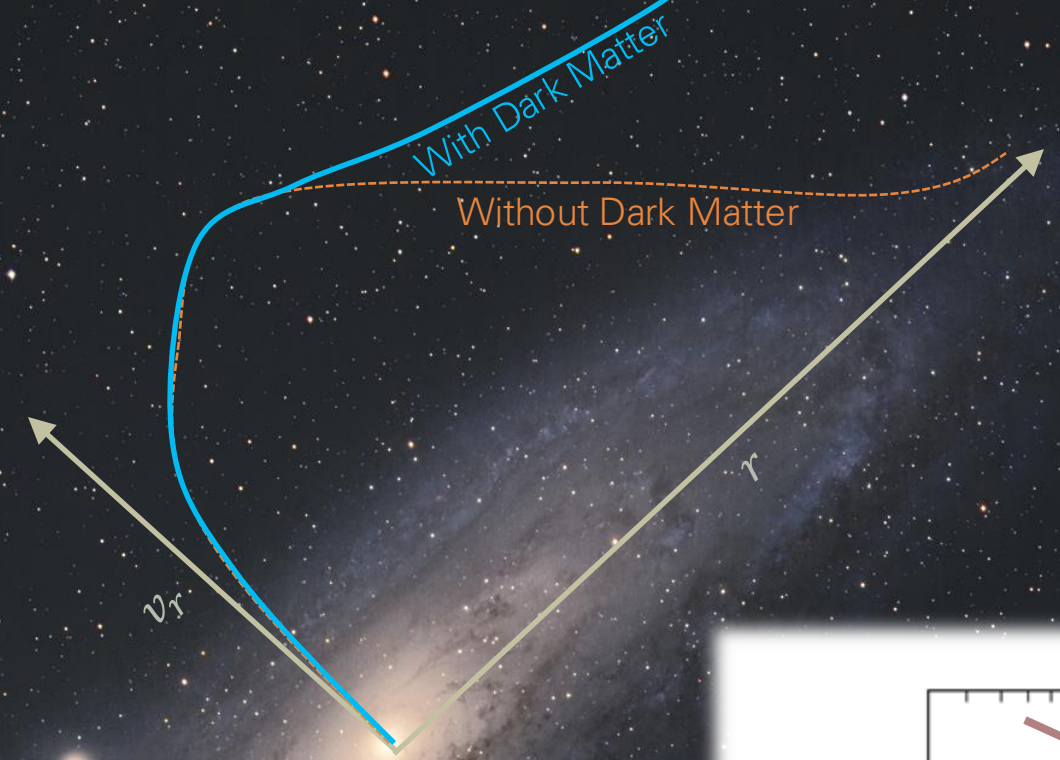
Rubin and Ford (1970)



Lina Necib, MIT



Circular Velocity → Gravitational Potential → Dark Matter Density



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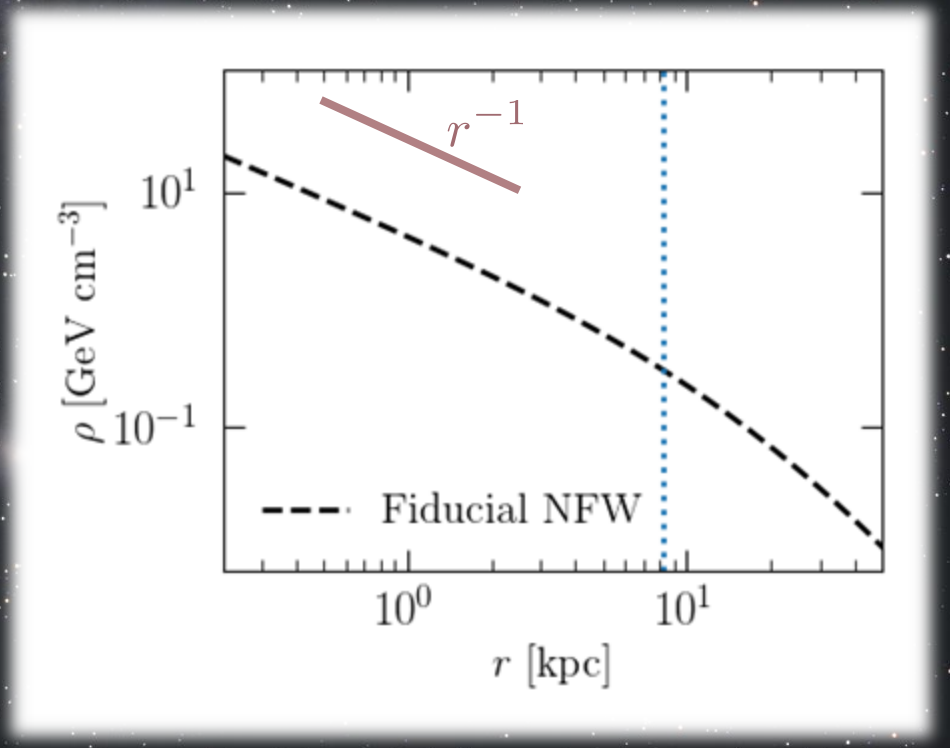
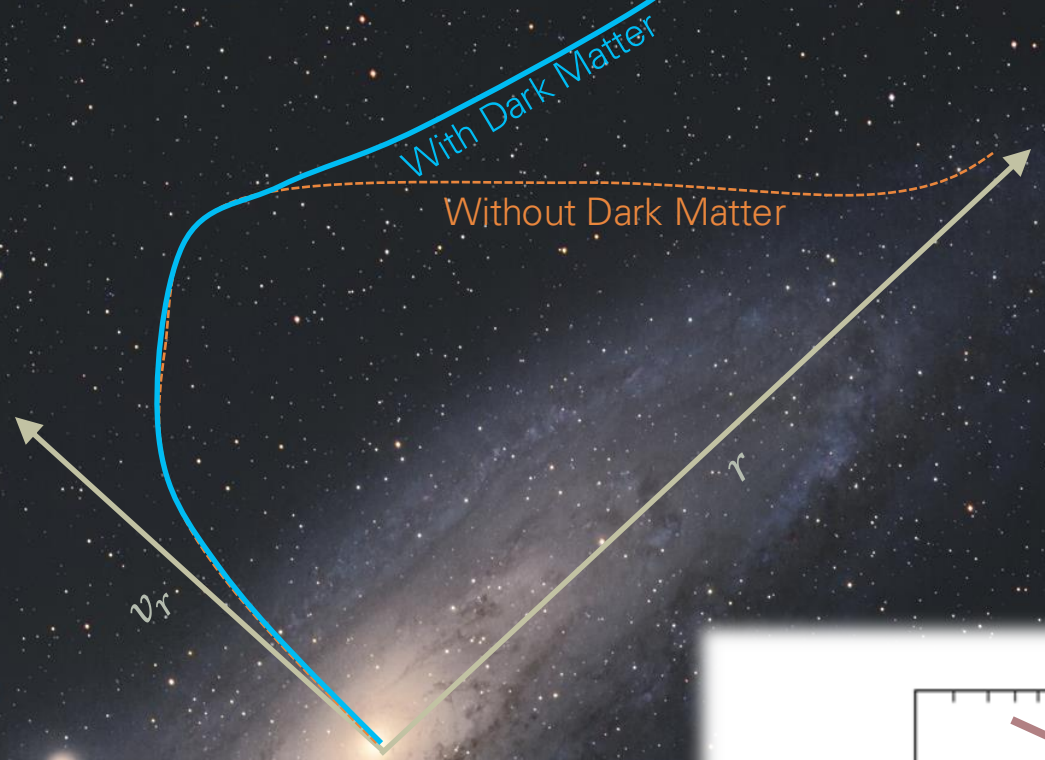


Navarro et al. (1997)



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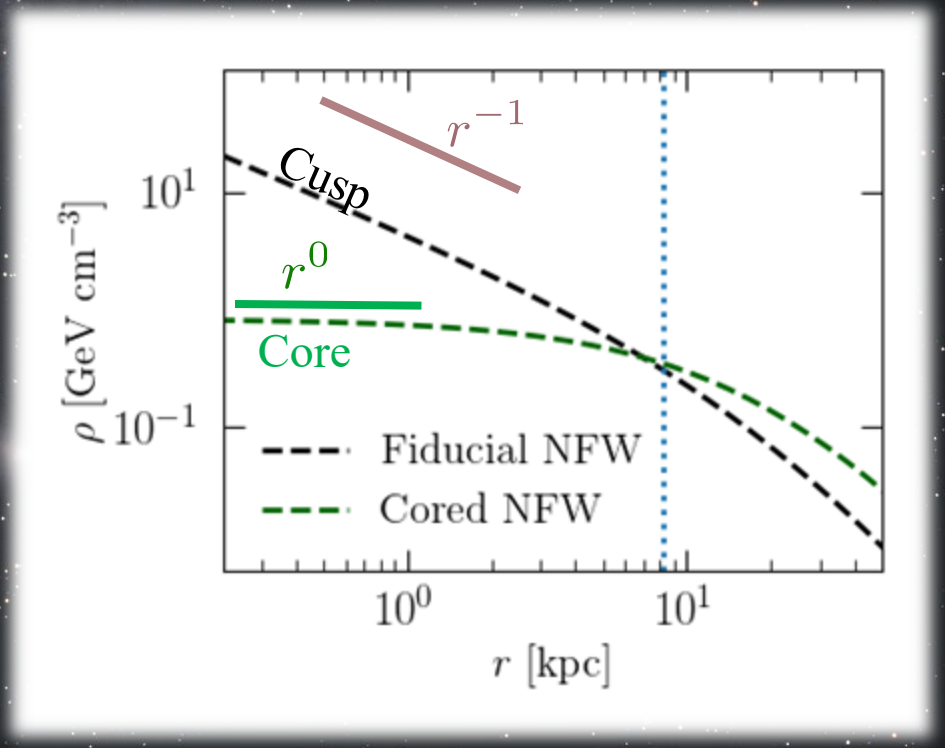
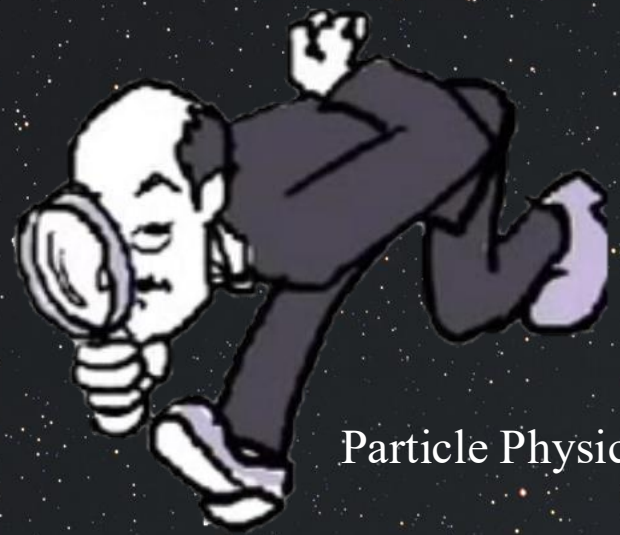
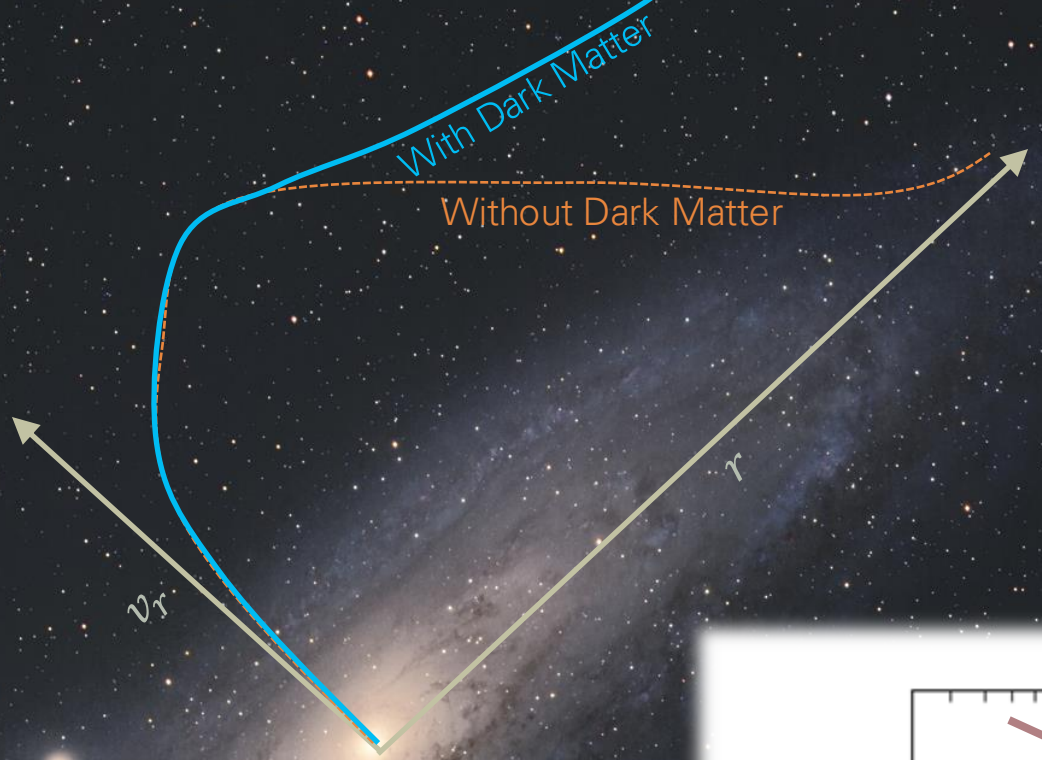




Particle Physics



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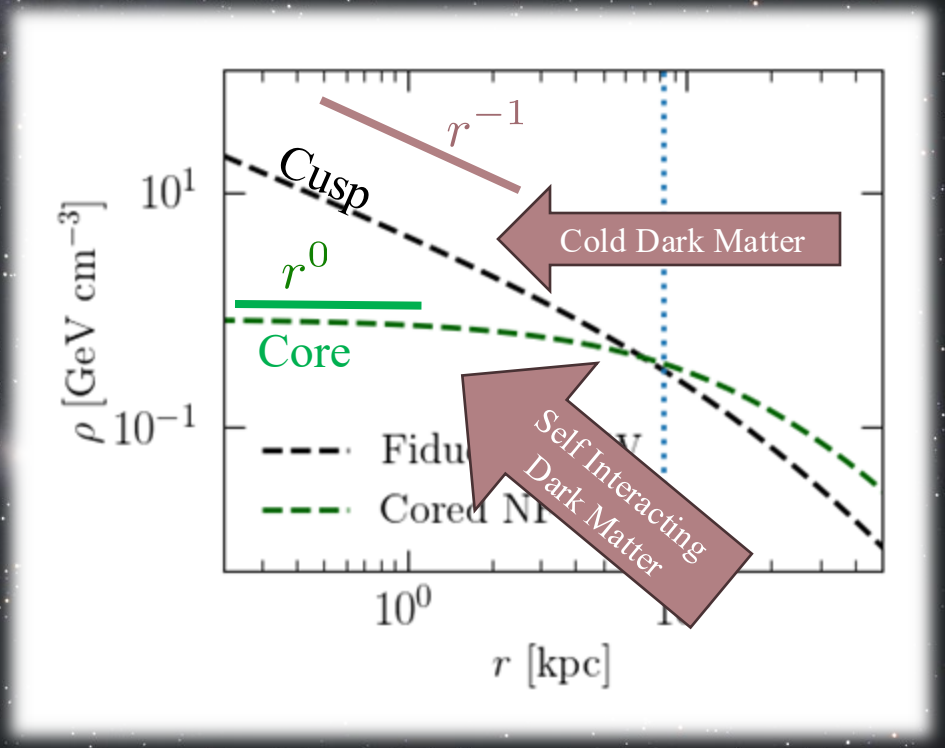
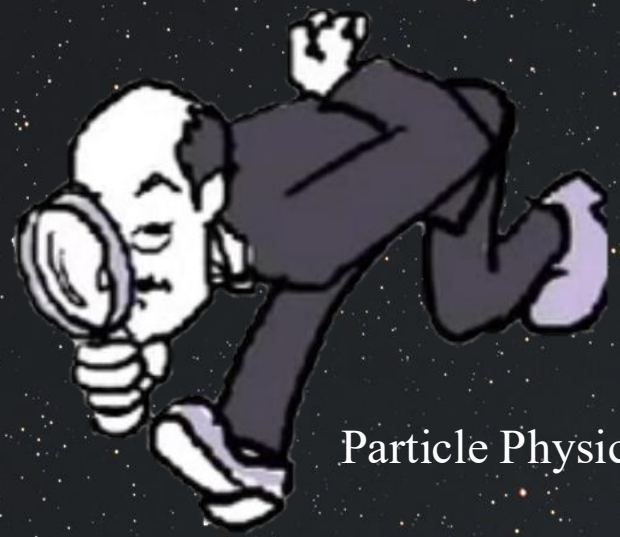
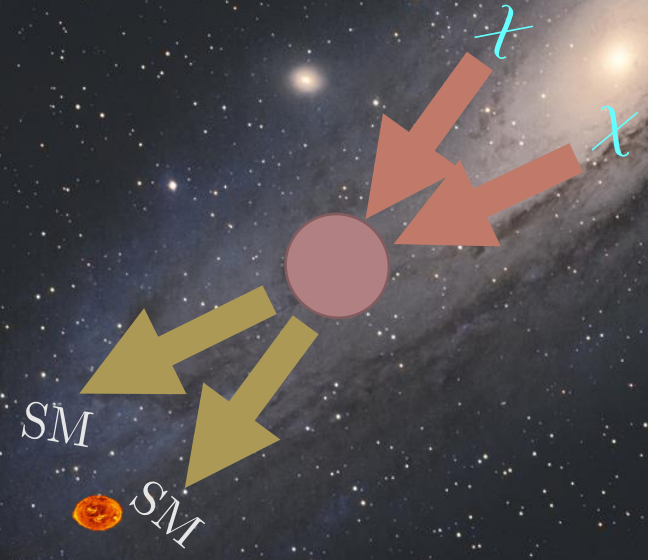
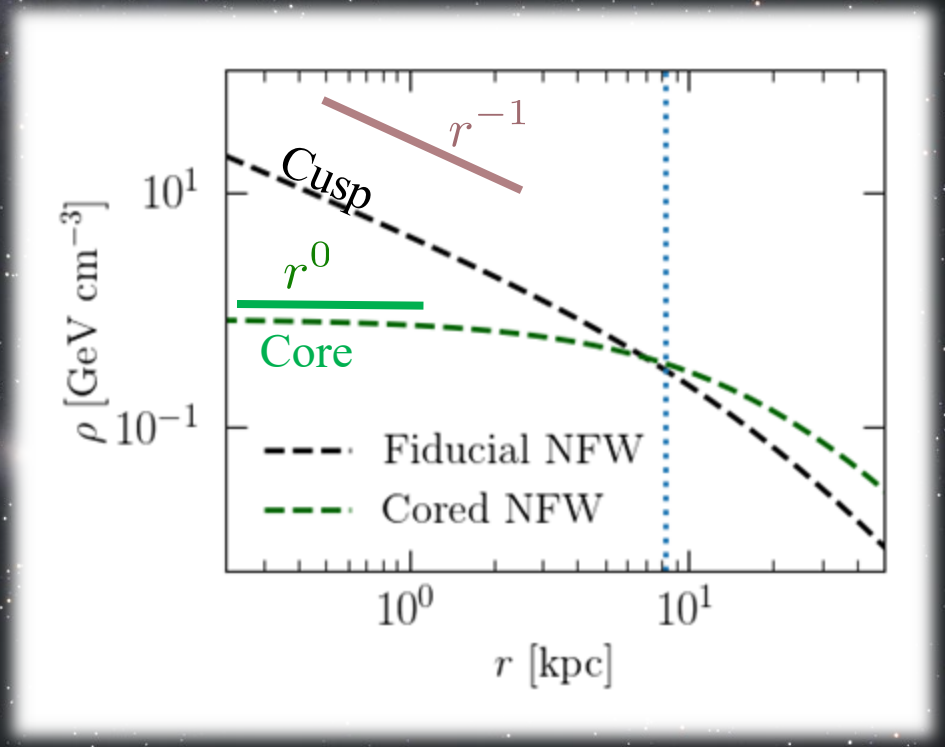


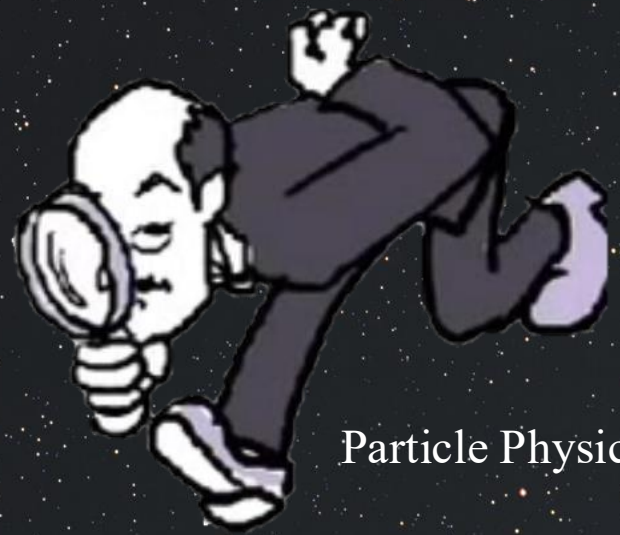


Particle Physics

Dark Matter
Annihilation Signal



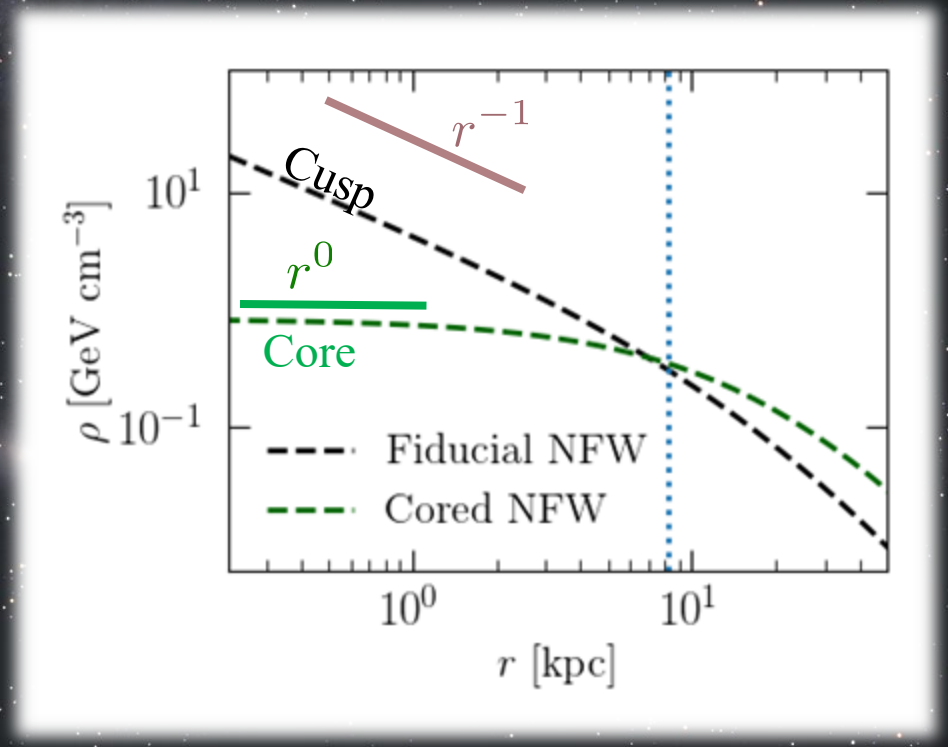


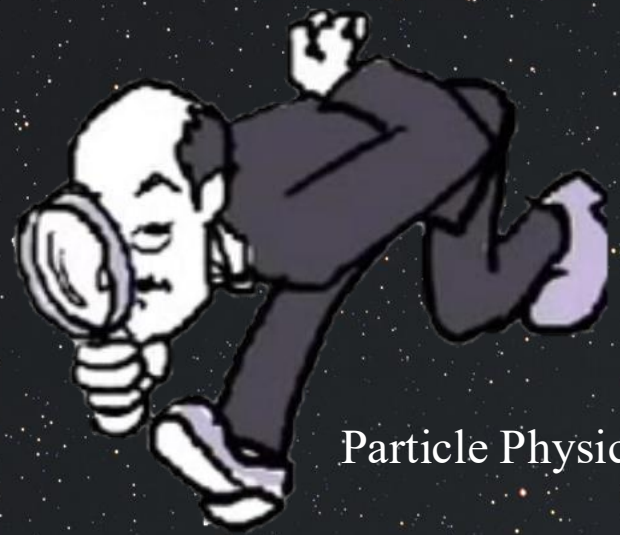


Particle Physics

Dark Matter
Annihilation Signal

$$\text{Signal} \propto \int \rho^2 \, \mathrm{ds}$$





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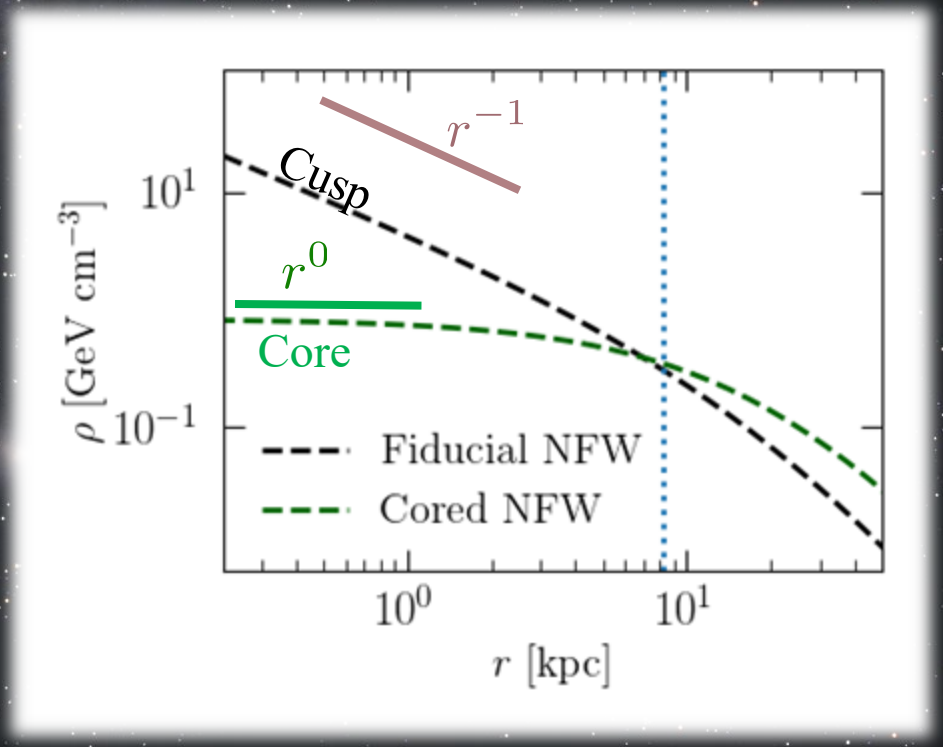
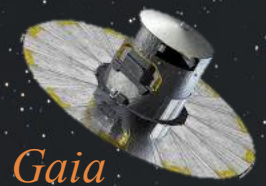
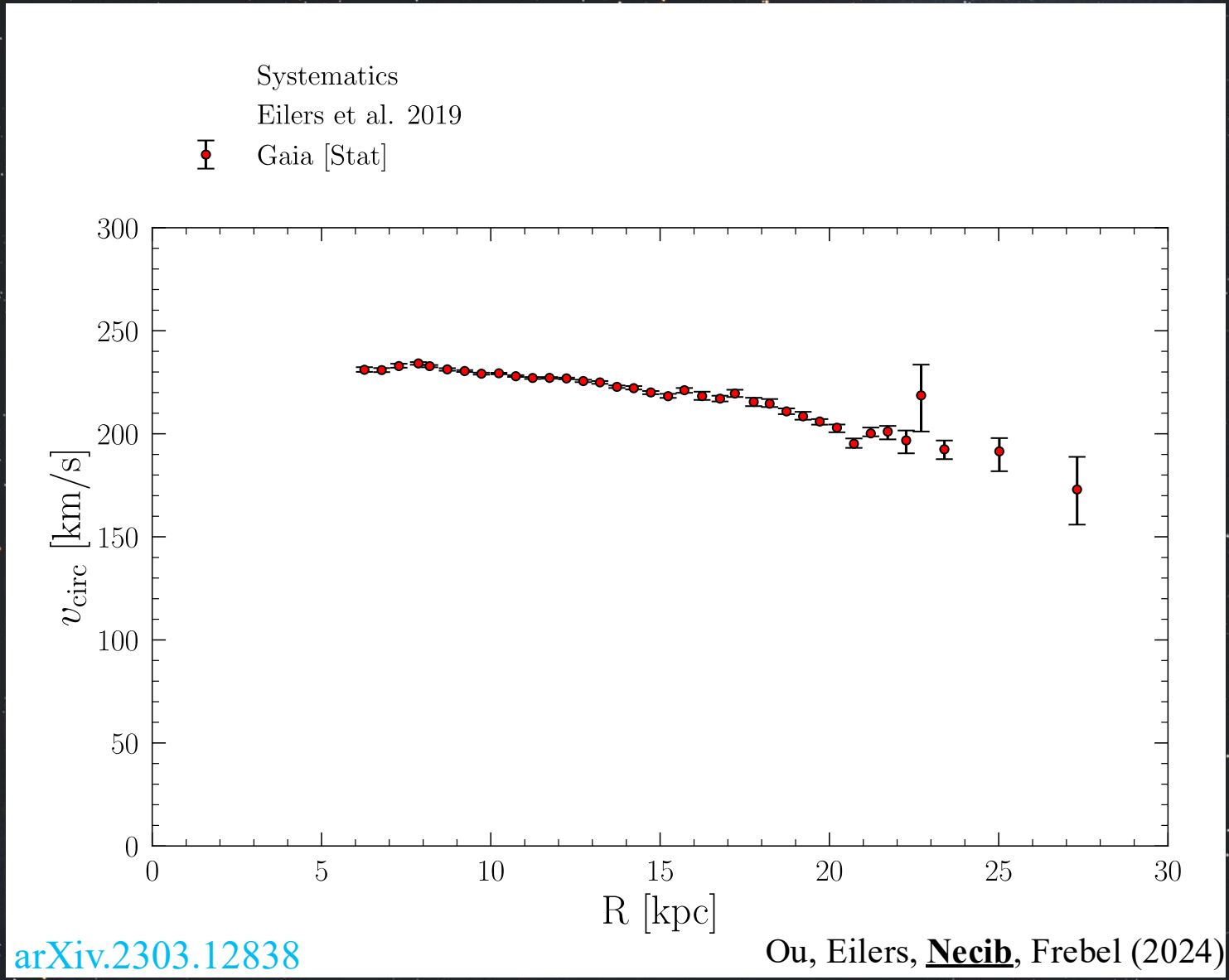


The effect can be orders of magnitude in signal expectation.

LET'S CONSTRUCT THE MILKY WAY'S CIRCULAR VELOCITY

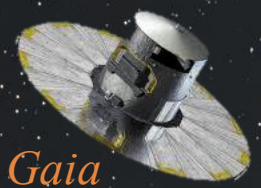


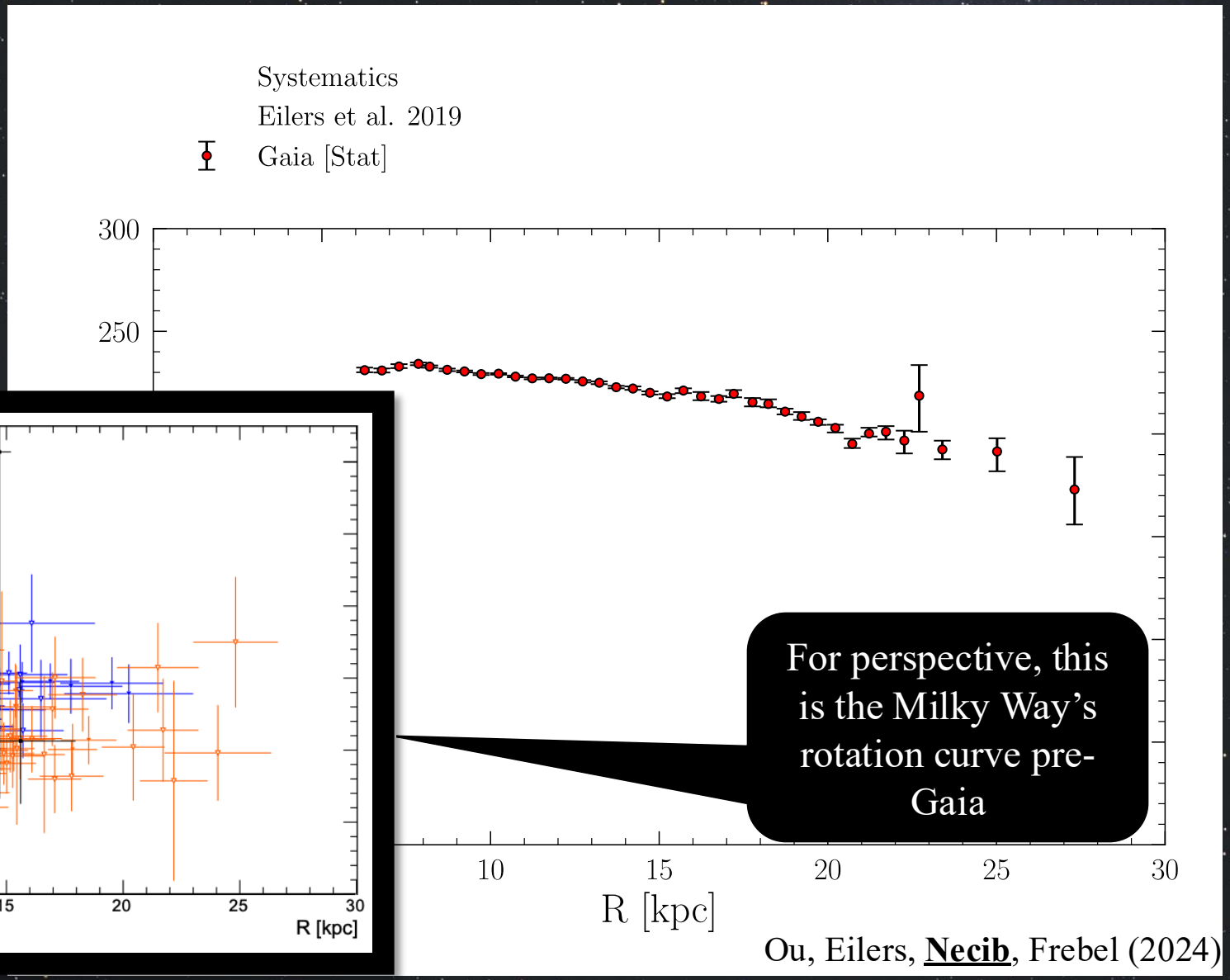


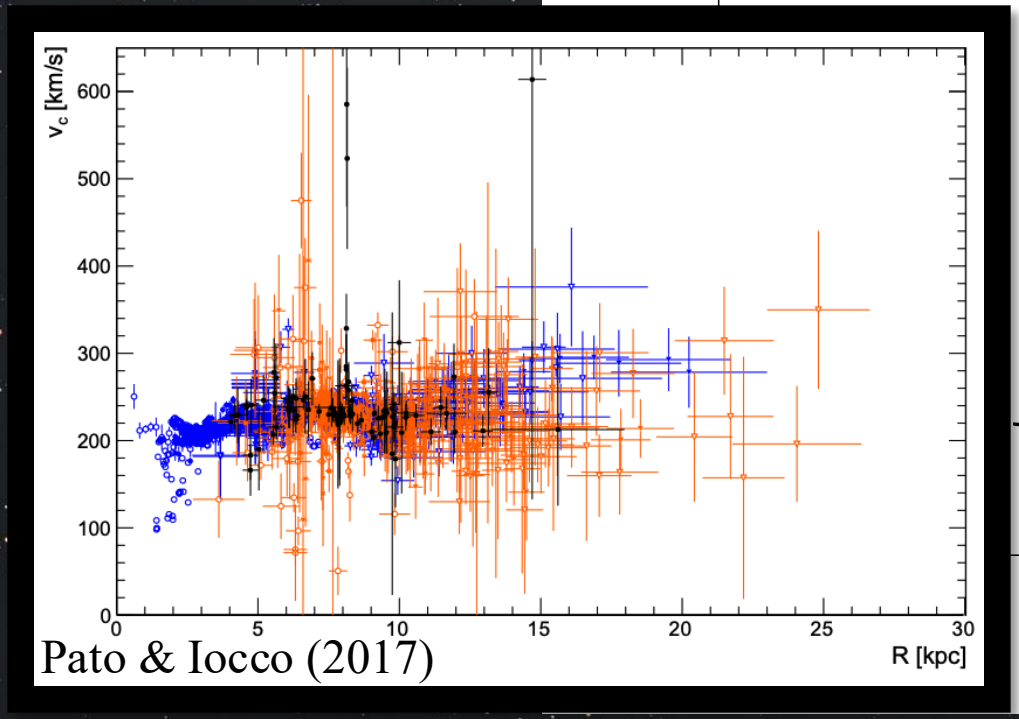
arXiv.2303.12838

Ou, Eilers, **Necib**, Frebel (2024)

LET'S CONSTRUCT THE MILKY WAY'S CIRCULAR VELOCITY







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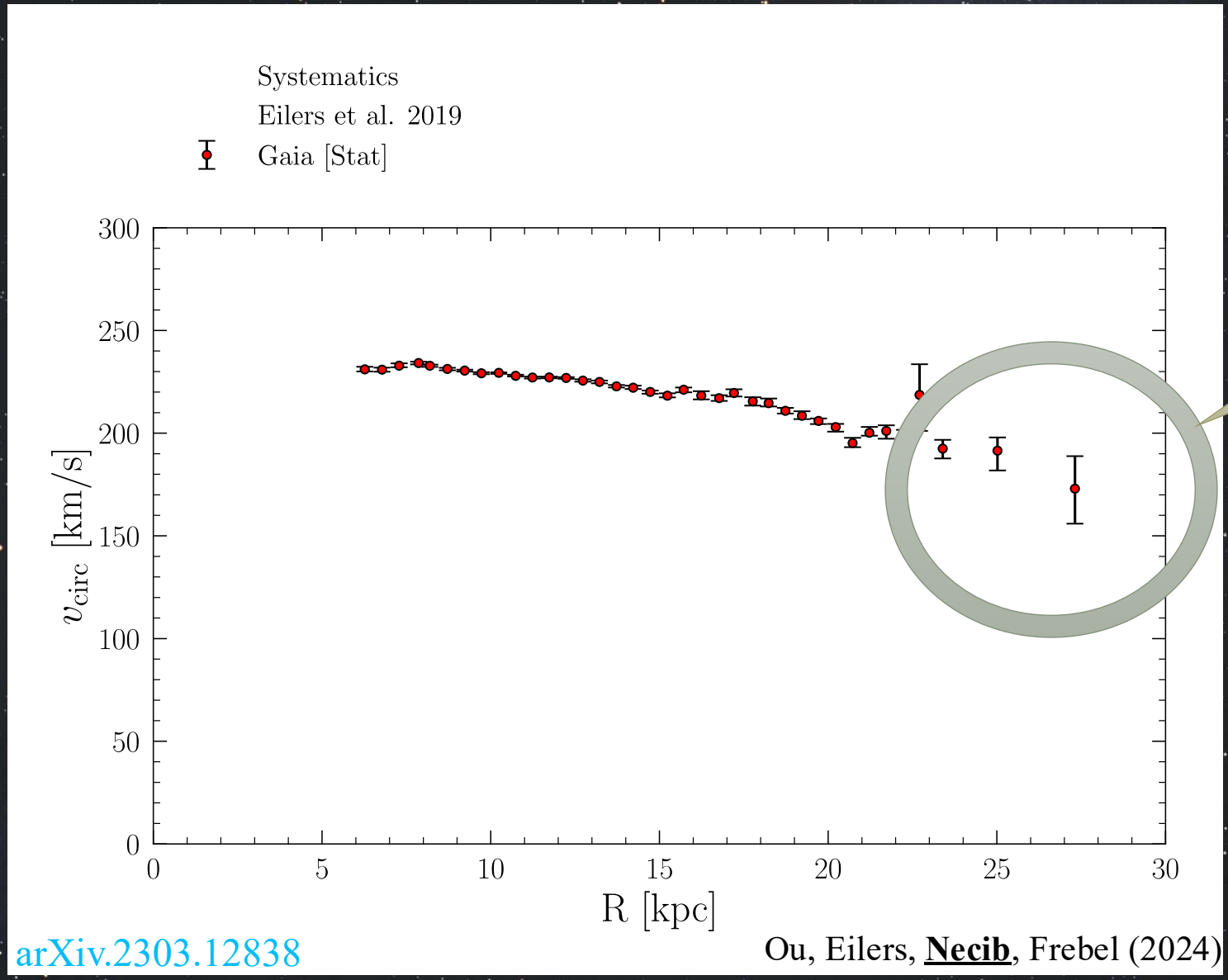

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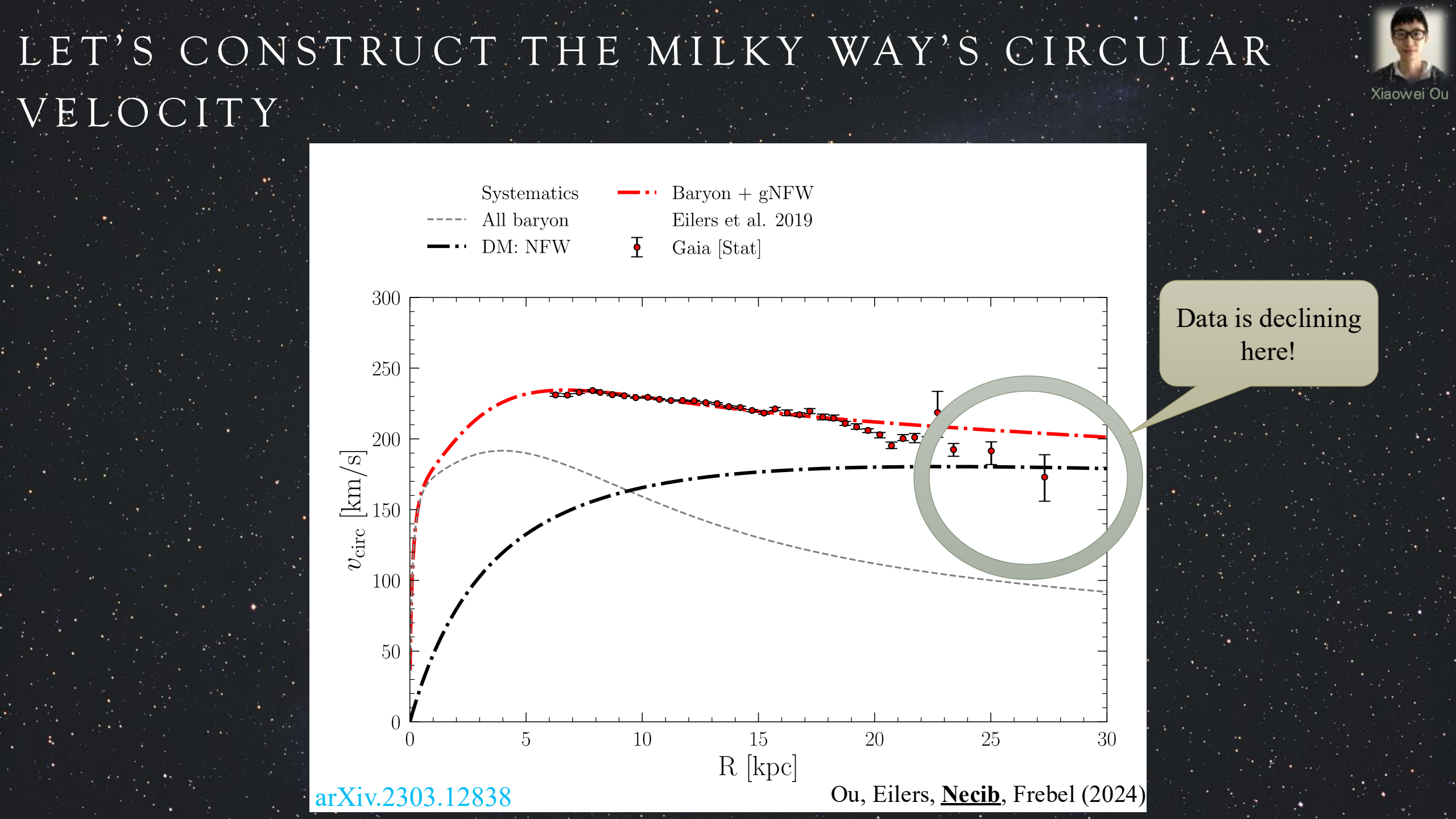

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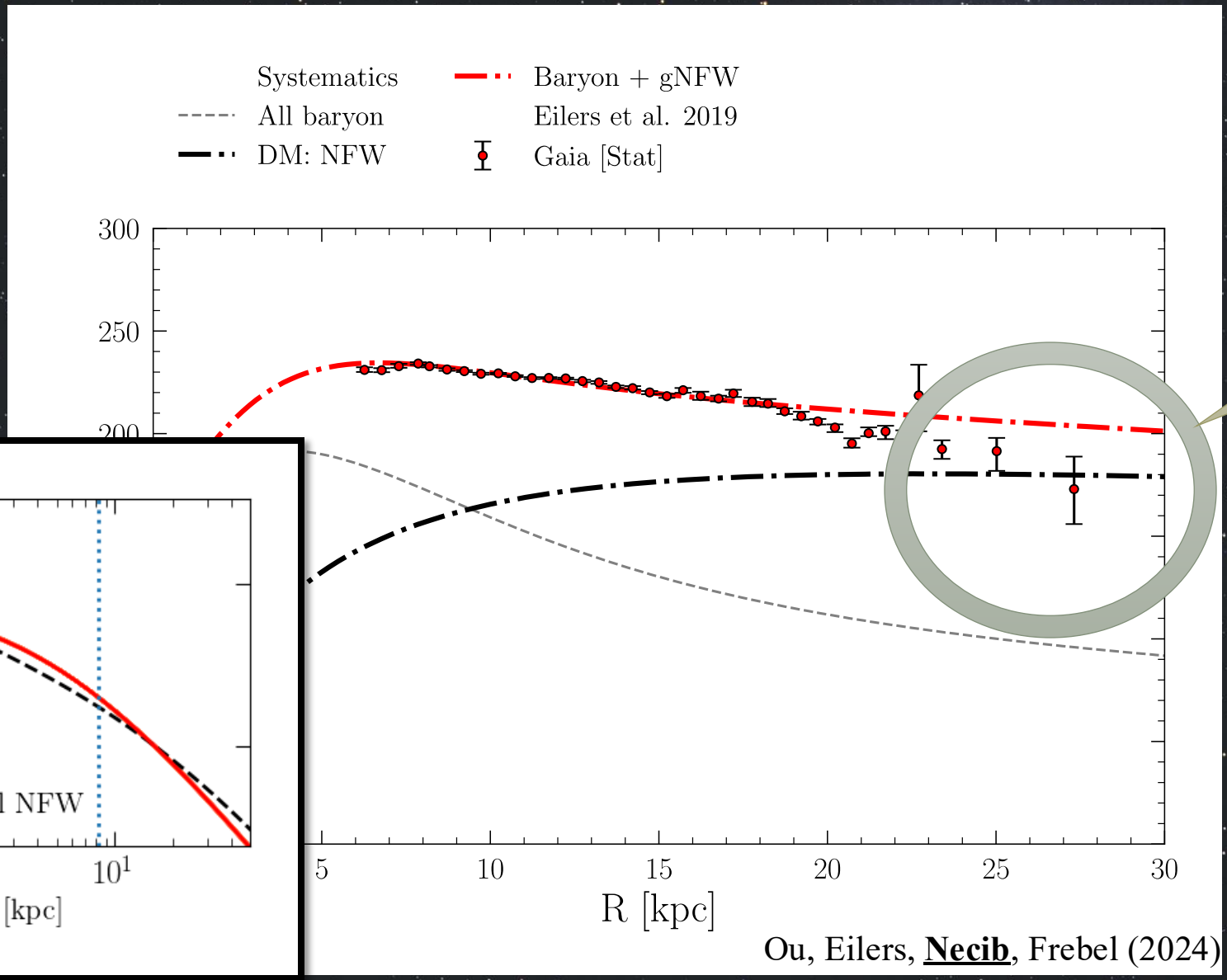
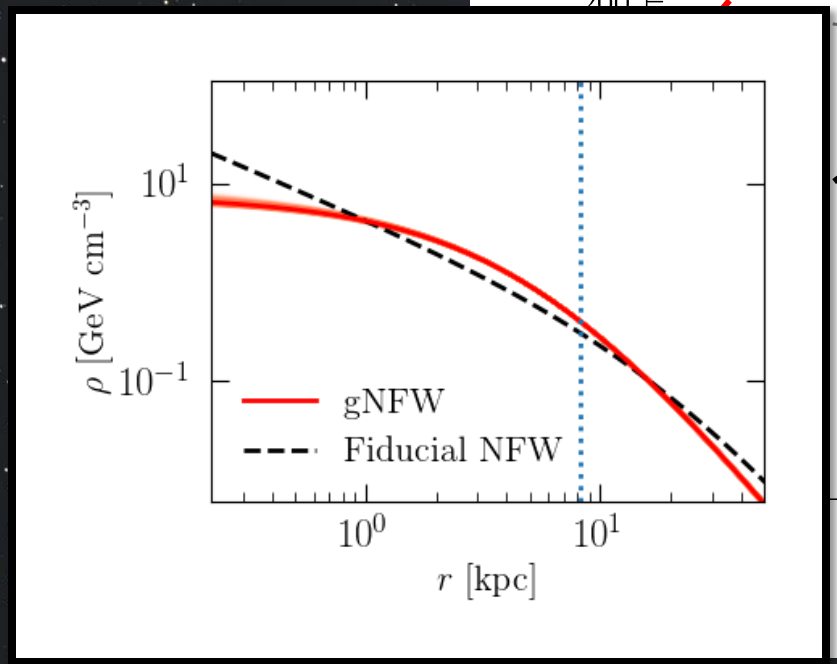
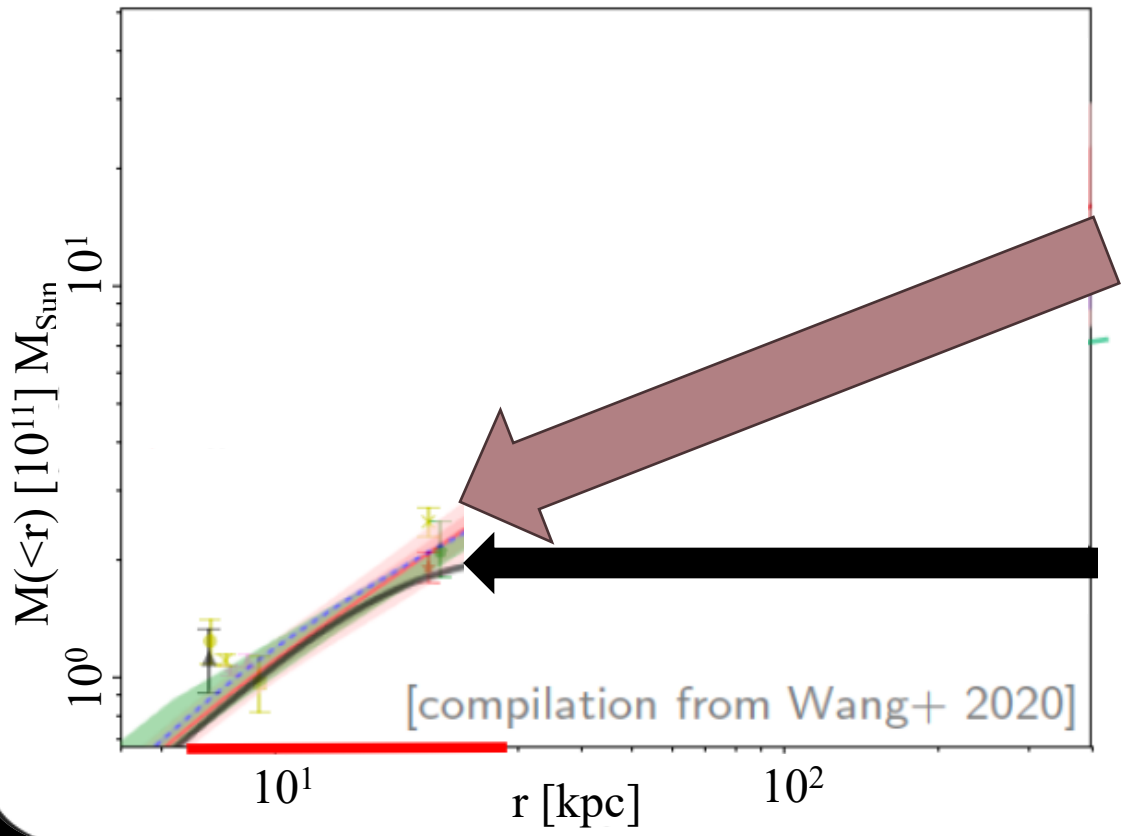




Ou, Eilers, **Necib**, Frebel (2024)

LETS PUT THAT CORED PROFILE IN CONTEXT.



GC+dSph (+LMC rewinding) [Correa Magnus & Vasiliev 2022]

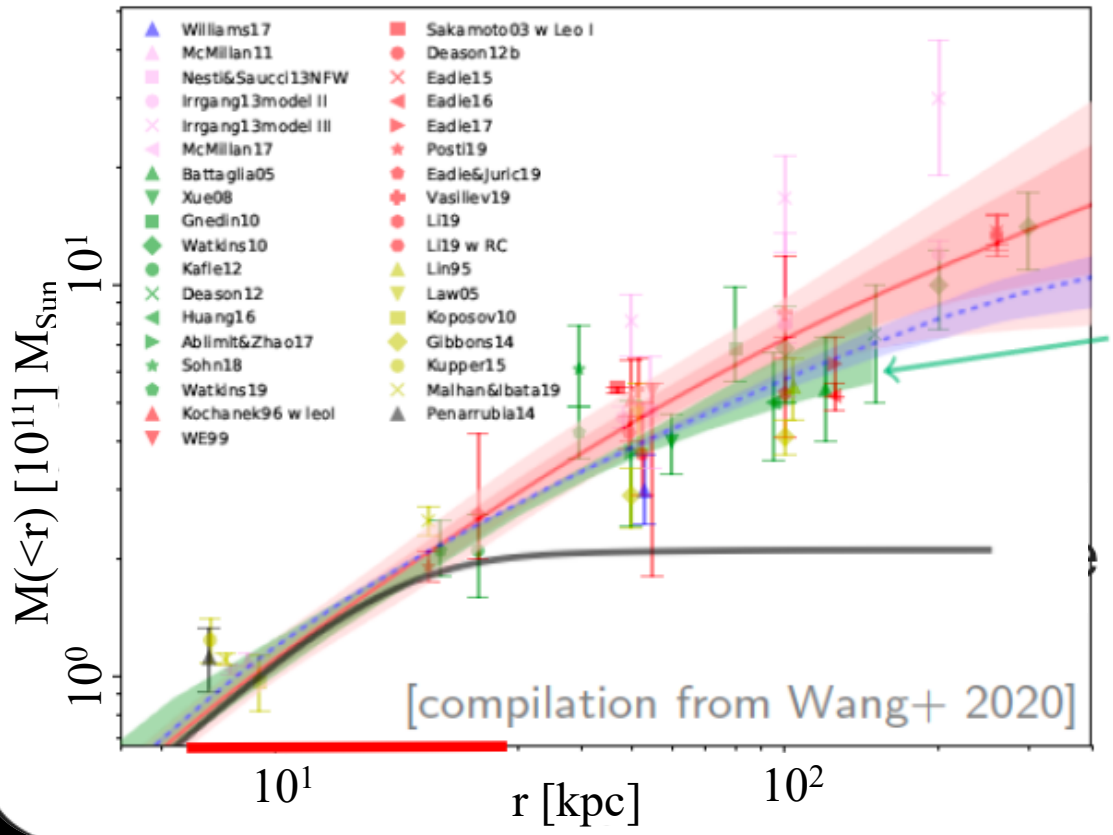
Sgr stream (incl. LMC) [Vasiliev+ 2021]

Orphan stream (incl. LMC) [Koposov+ 2022]

Best Fit from Ou, Eilers, Necib, Frebel (2024)

(Courtesy of Eugene Vasiliev)

SUCH CORED PROFILE IS IN TENSION WITH OTHER PROBES OF DARK MATTER IN THE MILKY WAY.



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LET'S REVIEW THE ASSUMPTION WE MADE

Observational Selection Function

Axisymmetry

Dynamical Equilibrium

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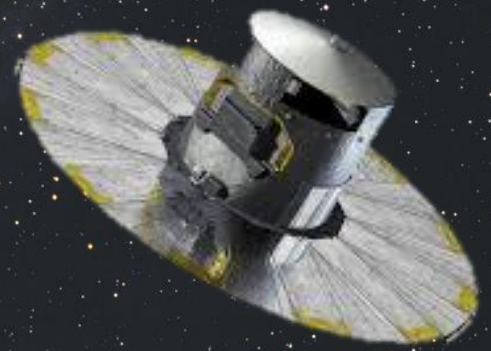
Dynamical Equilibrium



Is Gaia picking some stars over others?
Is that biasing the results of our circular velocity measurements?


Tri Nguyen


Xiaowei Ou

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Axisymmetry


Dynamical Equilibrium


Synthetic Gaia Survey (DR3) of a Simulated FIRE Galaxy

Is Gaia picking some stars over others?
Is that biasing the results of our circular velocity measurements?

Sanderson et al. (2020)
Nguyen, Ou, Panithanpaisal, Shipp, **Necib**, Sanderson, Wetzel (2024)

Publicly available at https://fire.northwestern.edu/ananke/


Tri Nguyen

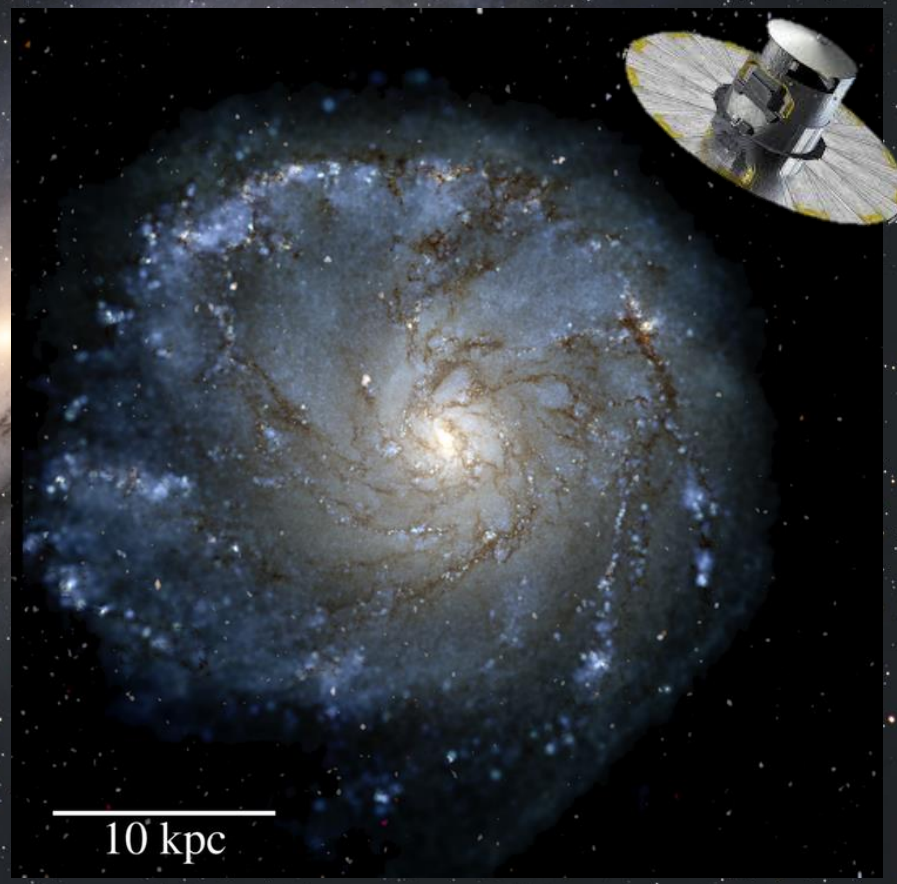

Xiaowei Ou

ARE THESE REASONABLE ASSUMPTIONS?


Observational Selection Function


Axisymmetry


Dynamical Equilibrium



Produced synthetic surveys and found no effect of the observational selection function

Nguyen, Ou,…, **Necib**, et al. (2024)
Ou, **Necib**, et al. (2025)

Lina Necib, MIT

ARE THESE REASONABLE ASSUMPTIONS?

Observational Selection Function

Axisymmetry

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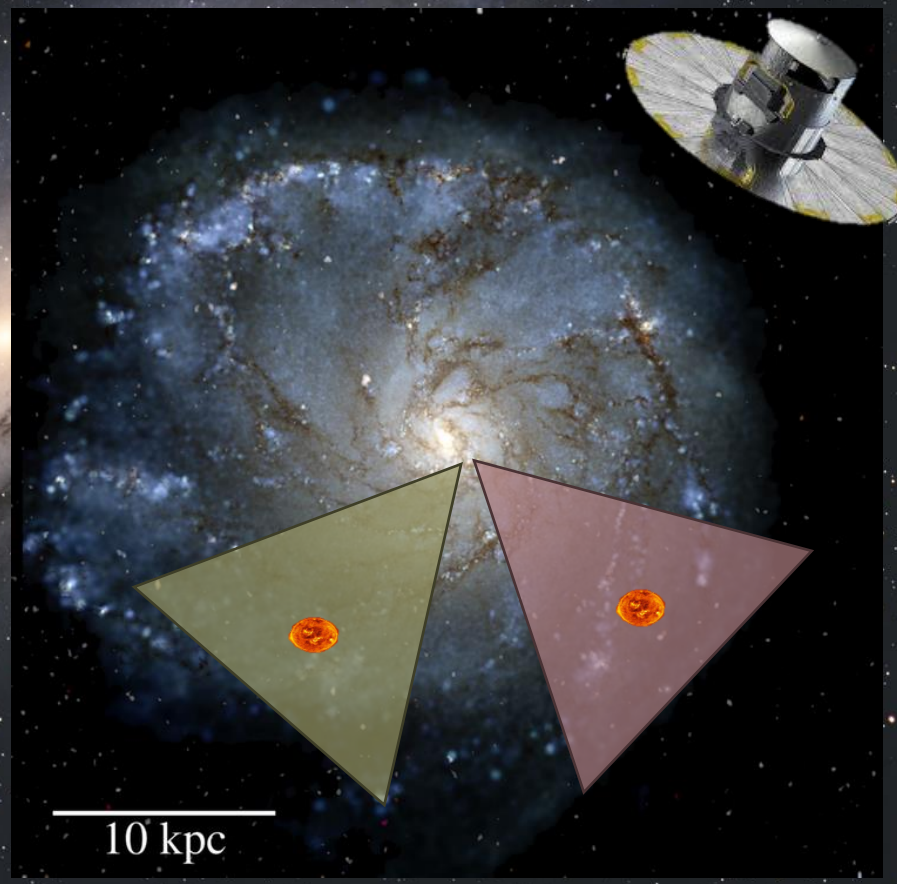


Nguyen, Ou,…, **Necib**, et al. (2024)
Ou, **Necib**, et al. (2025)

Lina Necib, MIT

Redid the analysis on different slices, and indeed there is a large effect

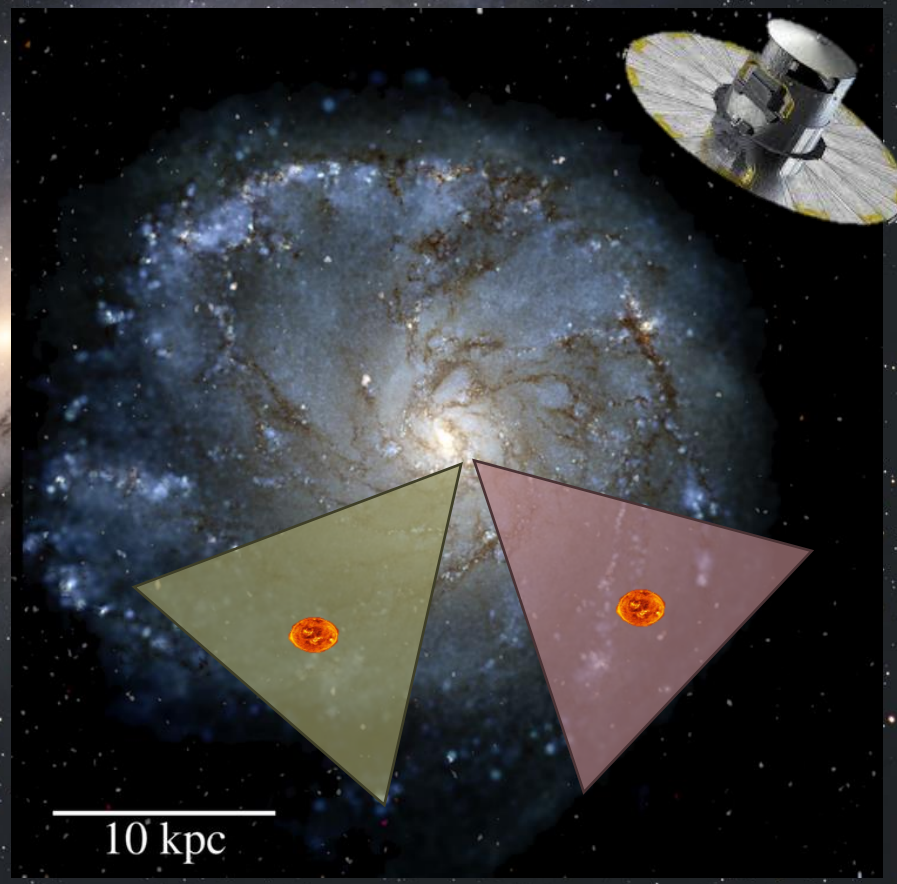

Xiaowei Ou

ARE THESE REASONABLE ASSUMPTIONS?

✓ 
Observational Selection Function

✗ 
Axisymmetry


Dynamical Equilibrium



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ARE THESE REASONABLE ASSUMPTIONS?

✓ 

Observational Selection Function

✗ 

Axisymmetry

Dynamical Equilibrium

Belokurov et al. (2018), Helmi et al. (Nature, 2018)
Necib et al. (Nature Astronomy 2020)
Nguyen, Ou,…, **Necib**, et al. (2024)
Ou, **Necib**, et al. (2025)

Lina Necib, MIT

We know that the Galaxy is not in equilibrium as evidenced by the different structures we found.


Xiaowei Ou

ARE THESE REASONABLE ASSUMPTIONS?

✓  Observational Selection Function

✗  Axisymmetry

✗  Dynamical Equilibrium

Belokurov et al. (2018), Helmi et al. (Nature, 2018)
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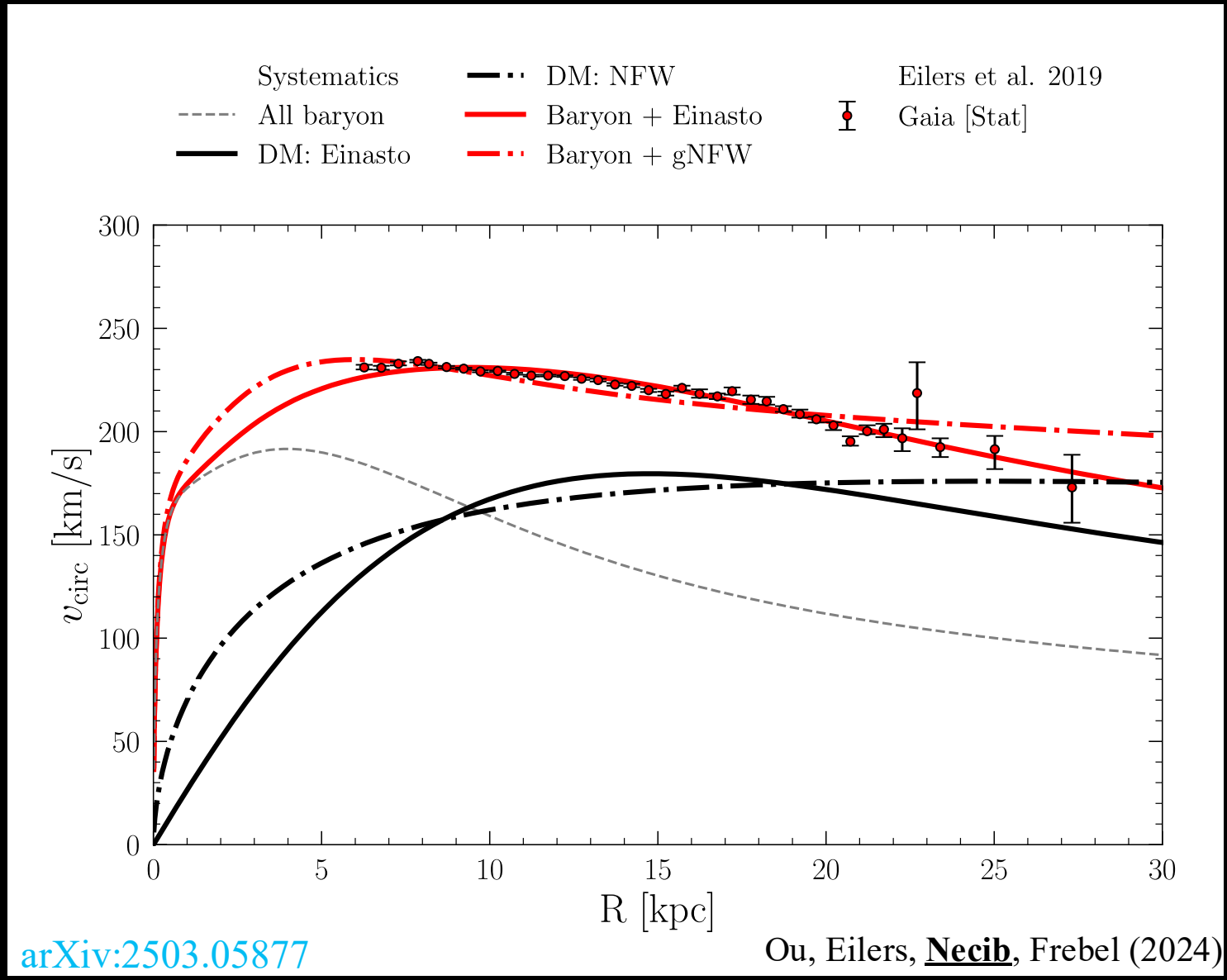
LET'S GO BACK TO THIS, WHAT ARE OUR REAL SYSTEMATICS?



arXiv:2503.05877

Ou, Eilers, **Necib**, Frebel (2024)

Ou, **Necib**, et al. (2025)

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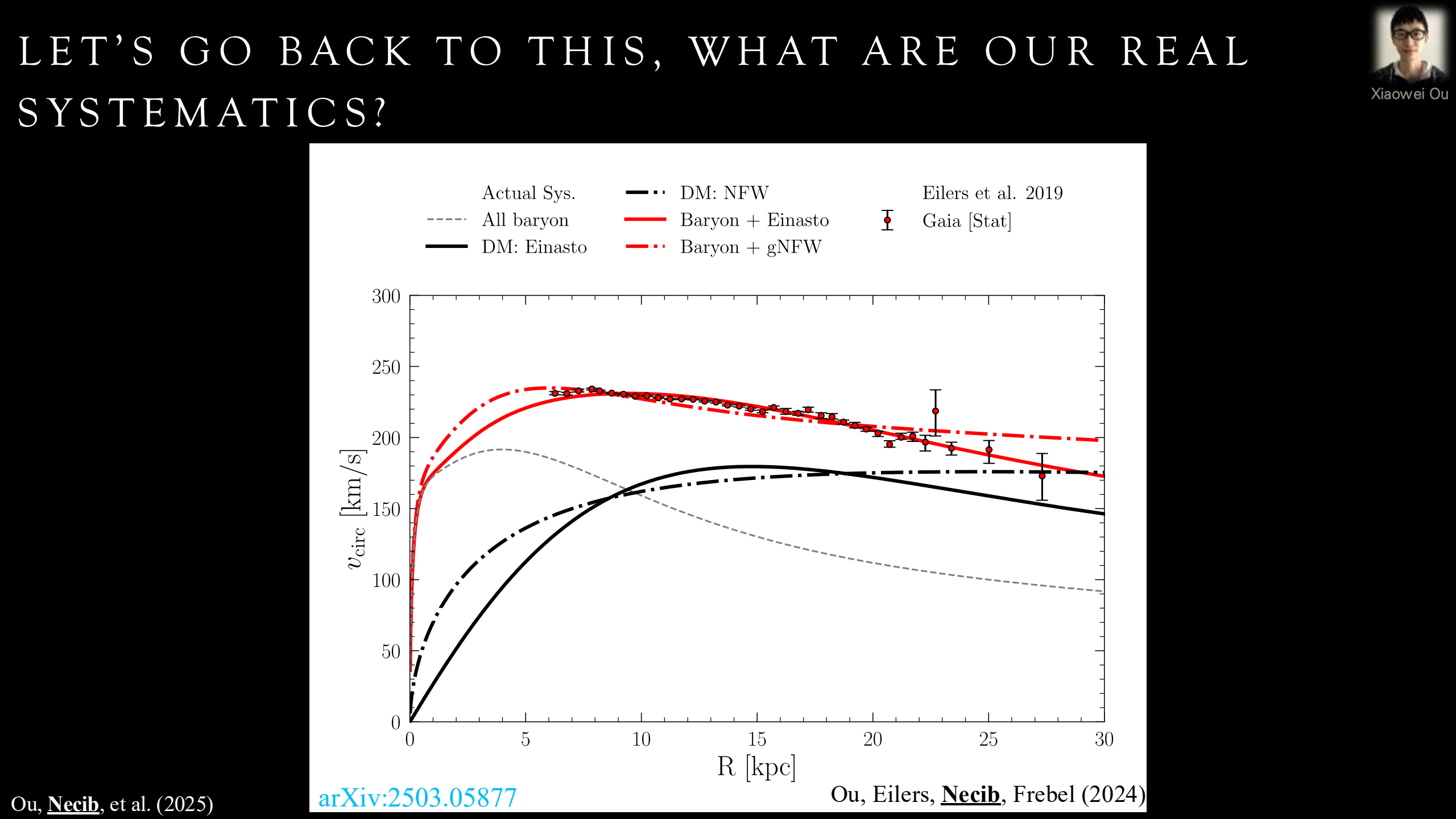

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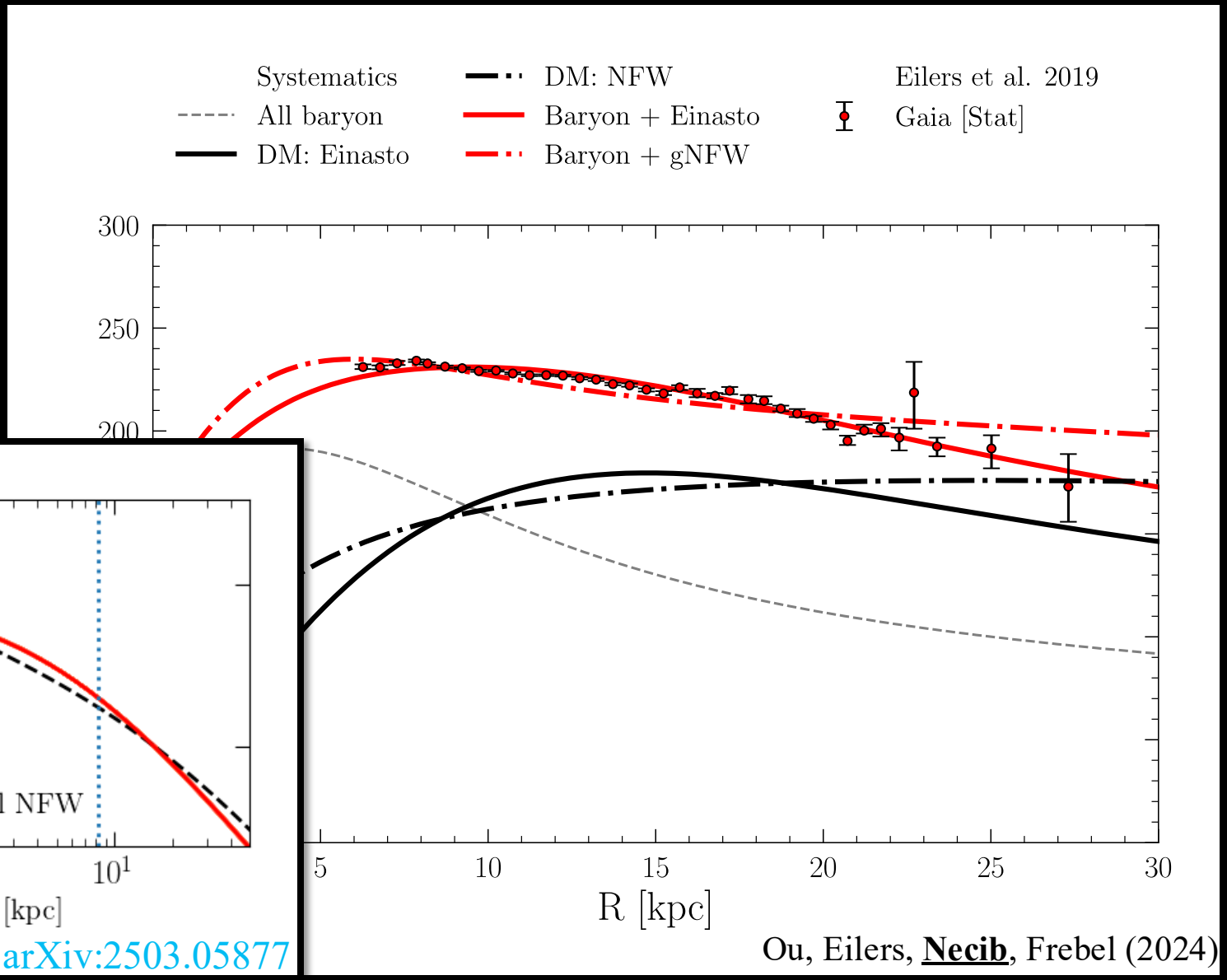
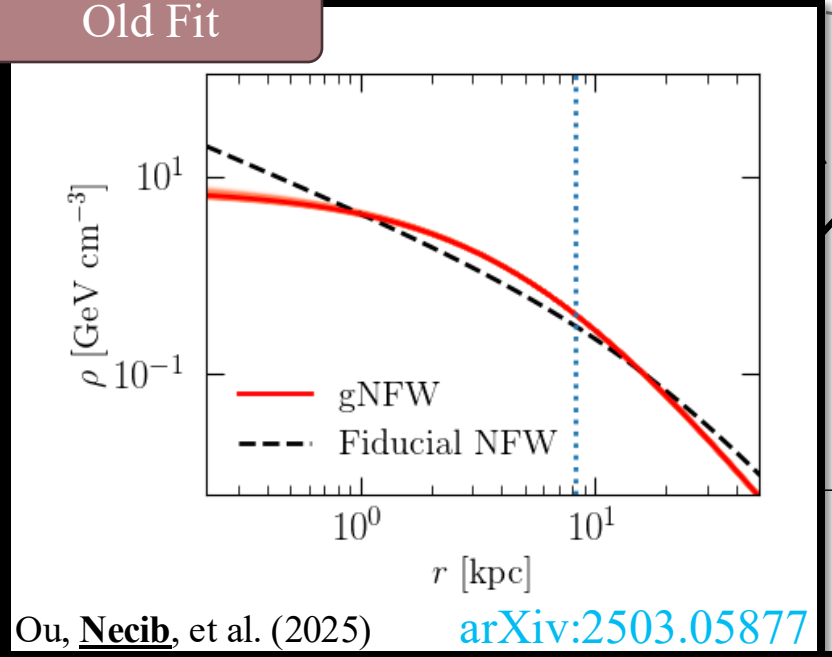


Ou, Eilers, **Necib**, Frebel (2024)

Old Fit



Ou, **Necib**, et al. (2025) arXiv:2503.05877

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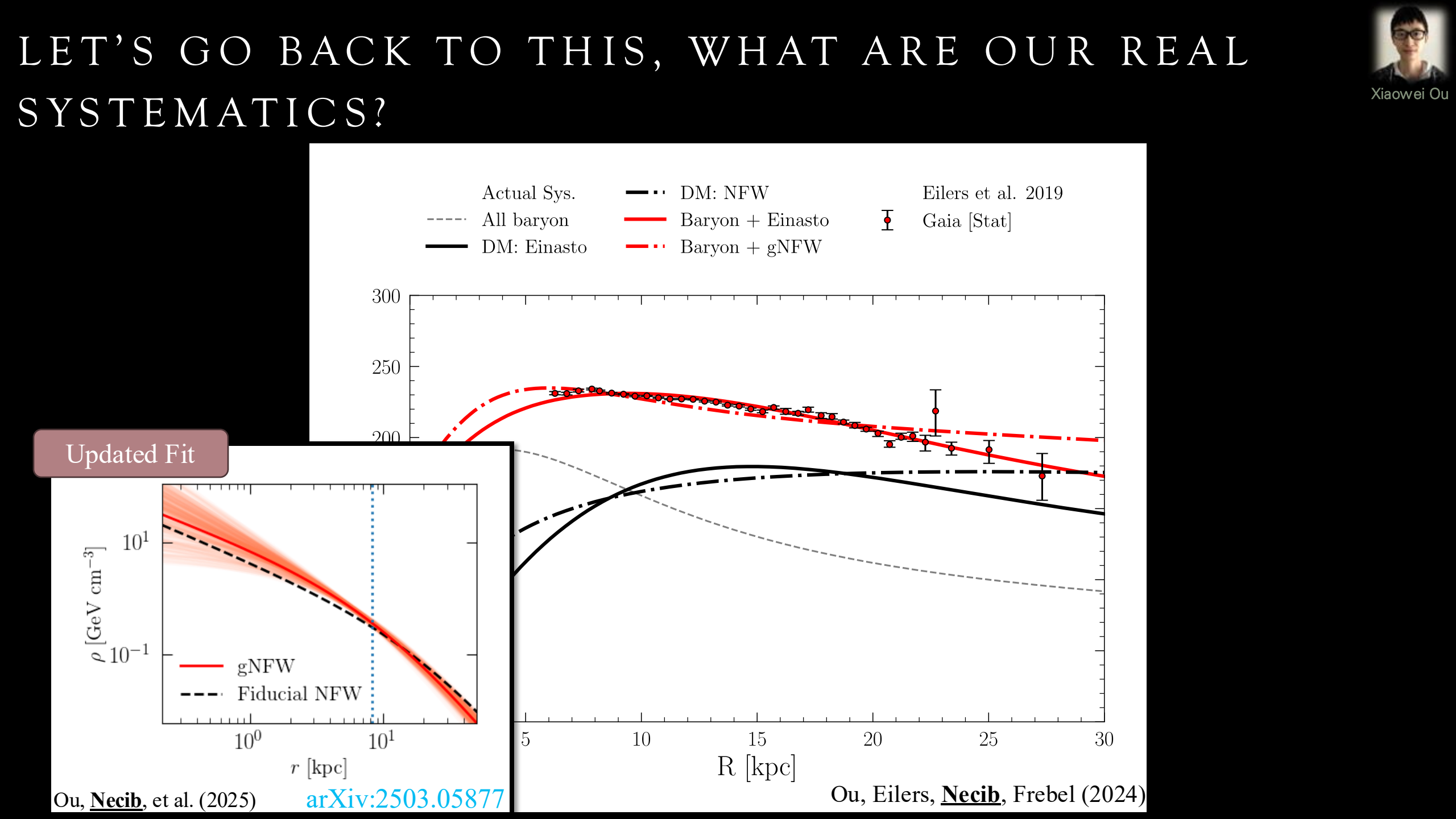

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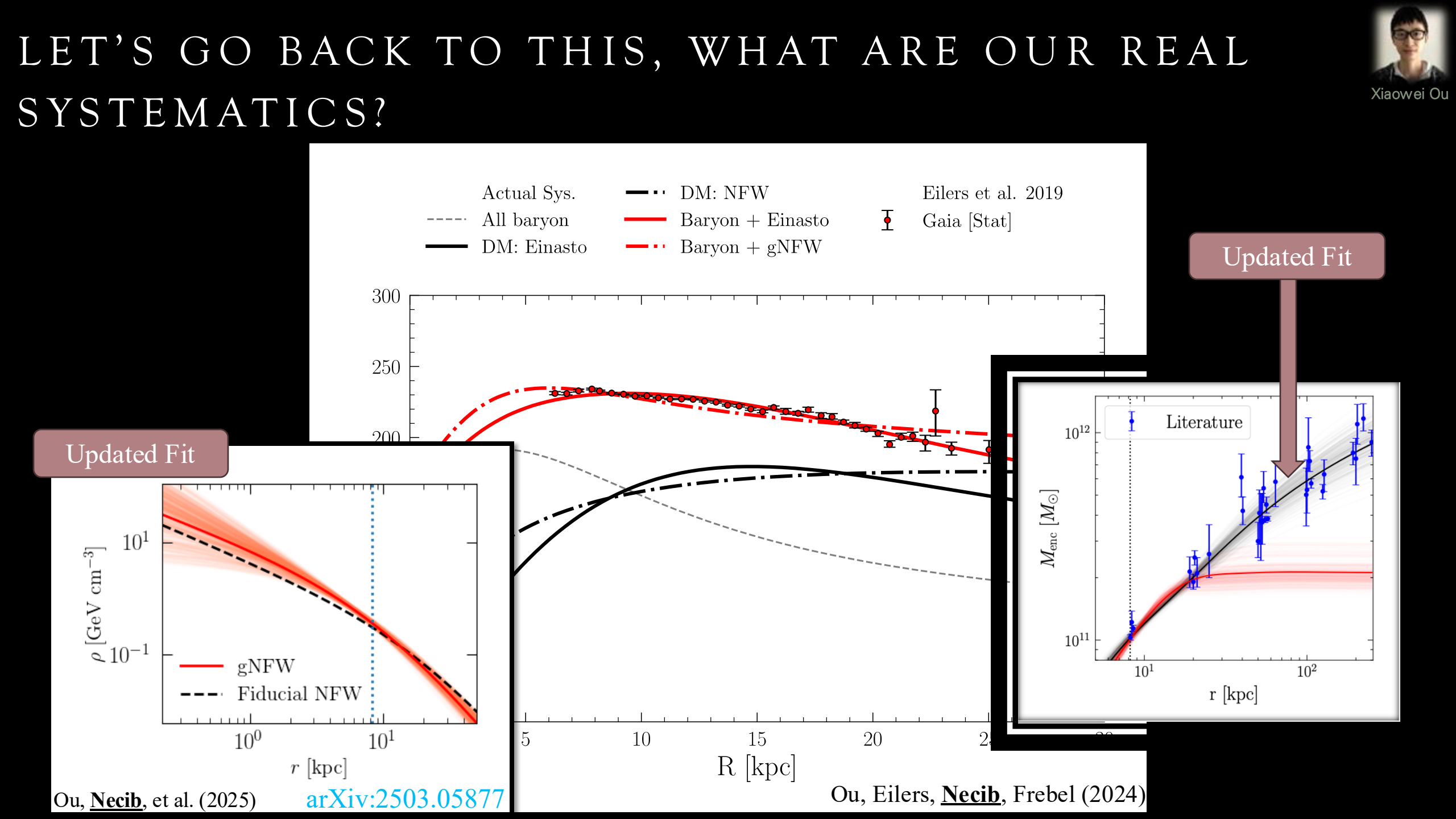
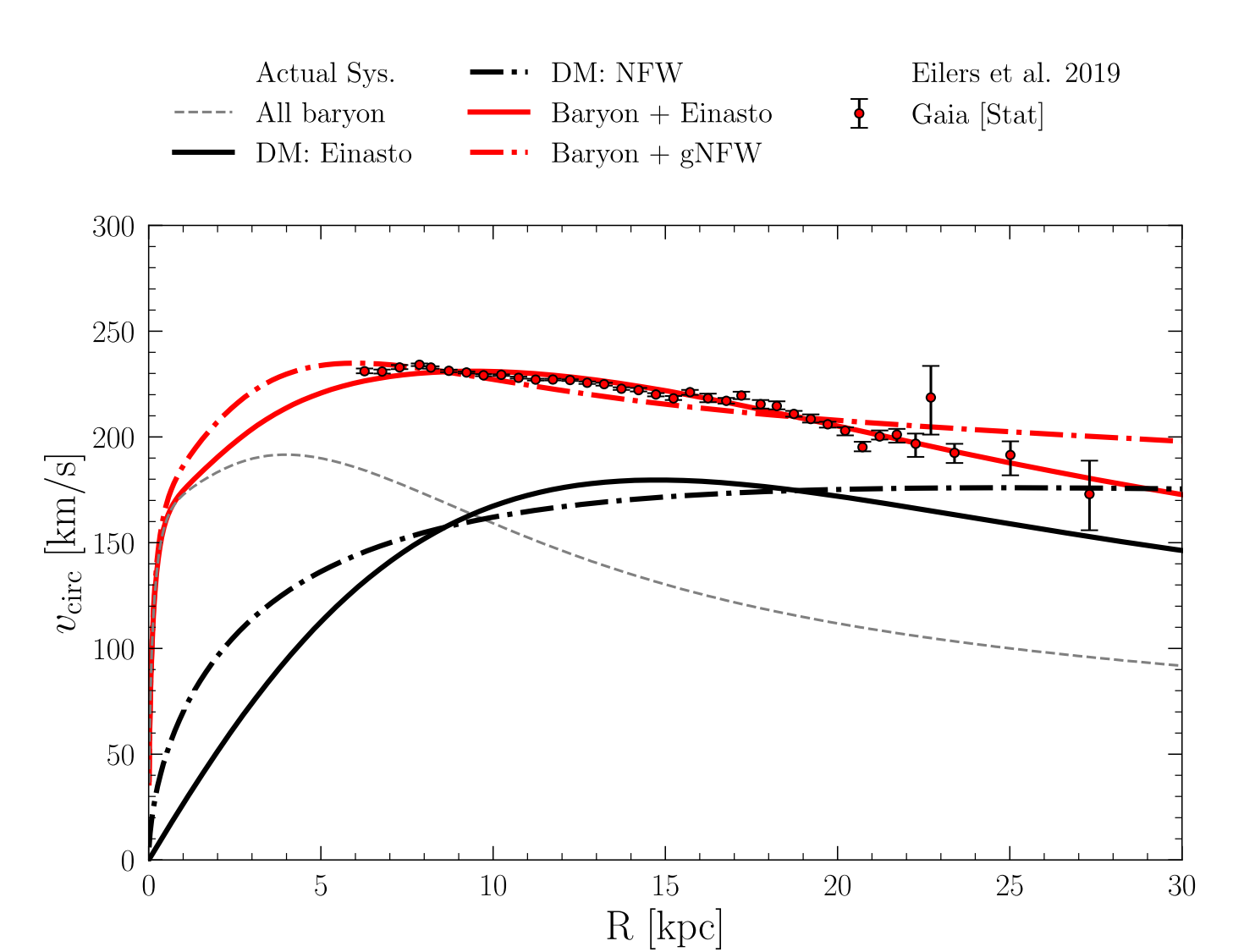


Ou, **Necib**, et al. (2025) arXiv:2503.05877

Ou, Eilers, **Necib**, Frebel (2024)

WE HAVE A BETTER FIT, THAT IS NOT IN TENSION, BUT WE STILL NEED TO SOLVE DARK MATTER...



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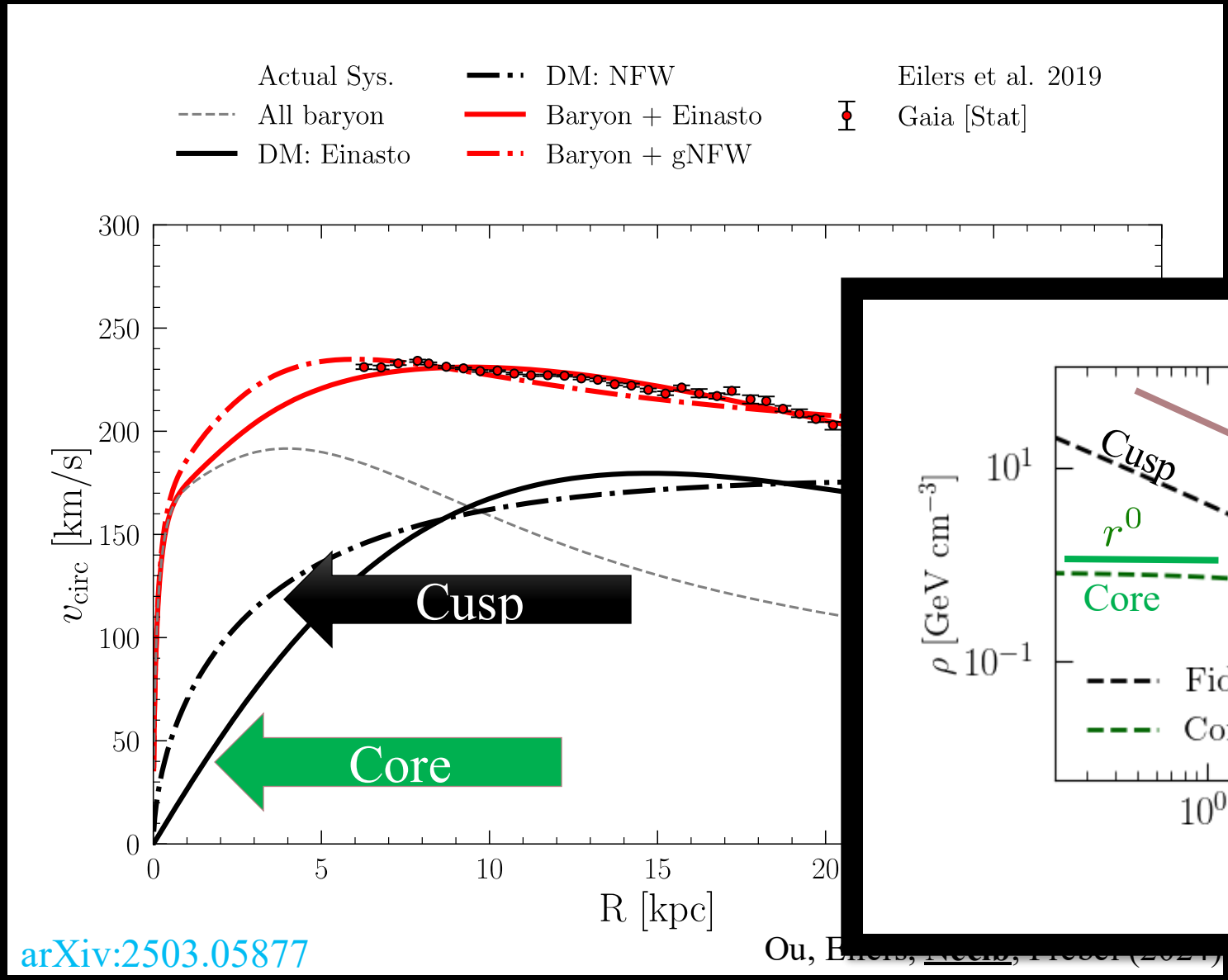
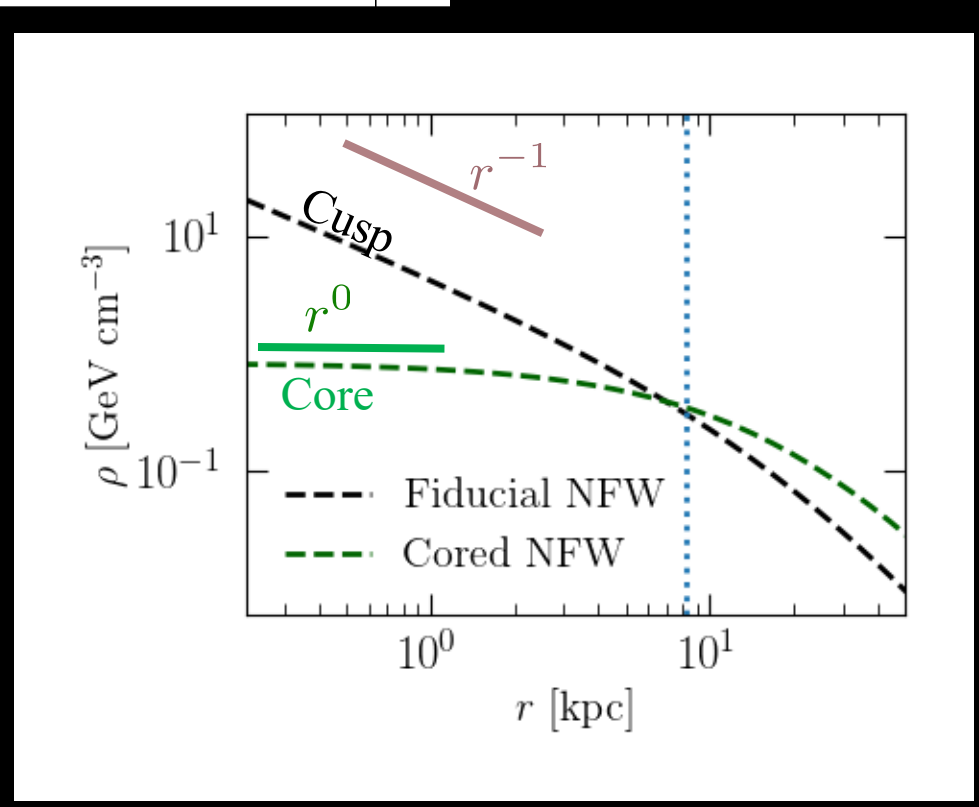
arXiv:2503.05877

Ou, Eilers, **Necib**, Frebel (2024)

WE HAVE A BETTER FIT, THAT IS NOT IN TENSION, BUT WE STILL NEED TO SOLVE DARK MATTER...





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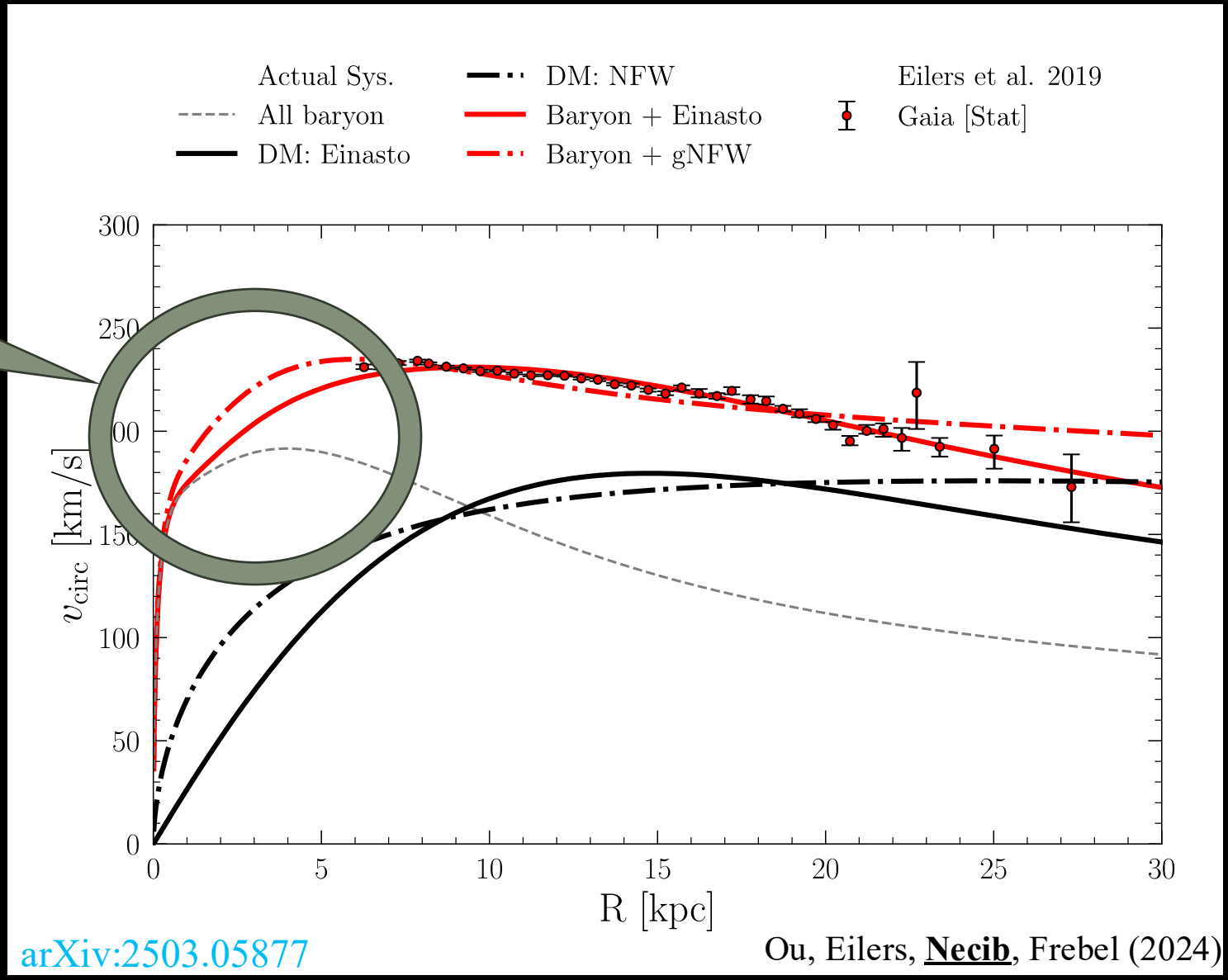


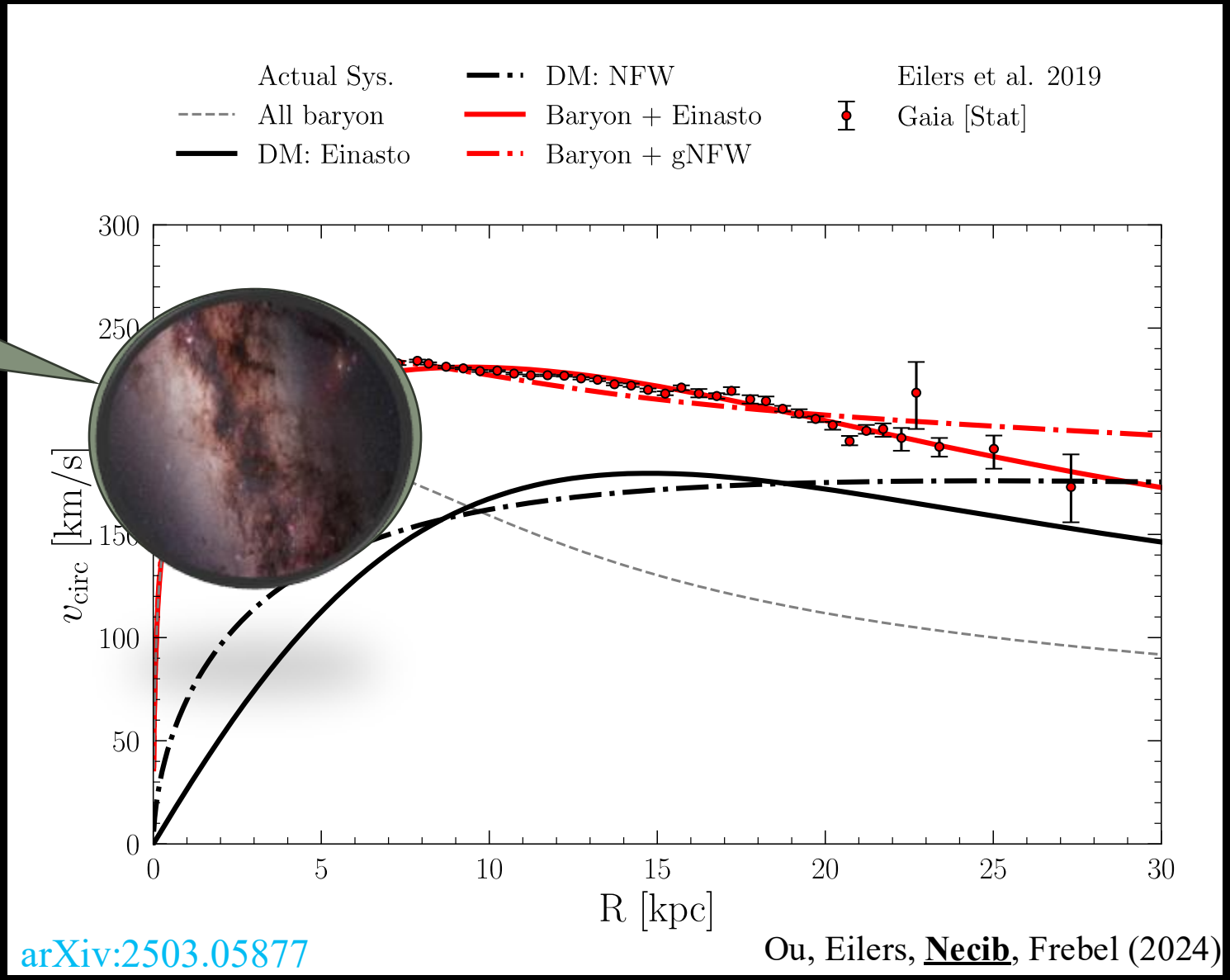
Ou, **Necib**, et al. (2025)

arXiv:2503.05877

Ou, Eilers, **Necib**, Frebel (2024)

BUT DUST...

What I really want, is data here!



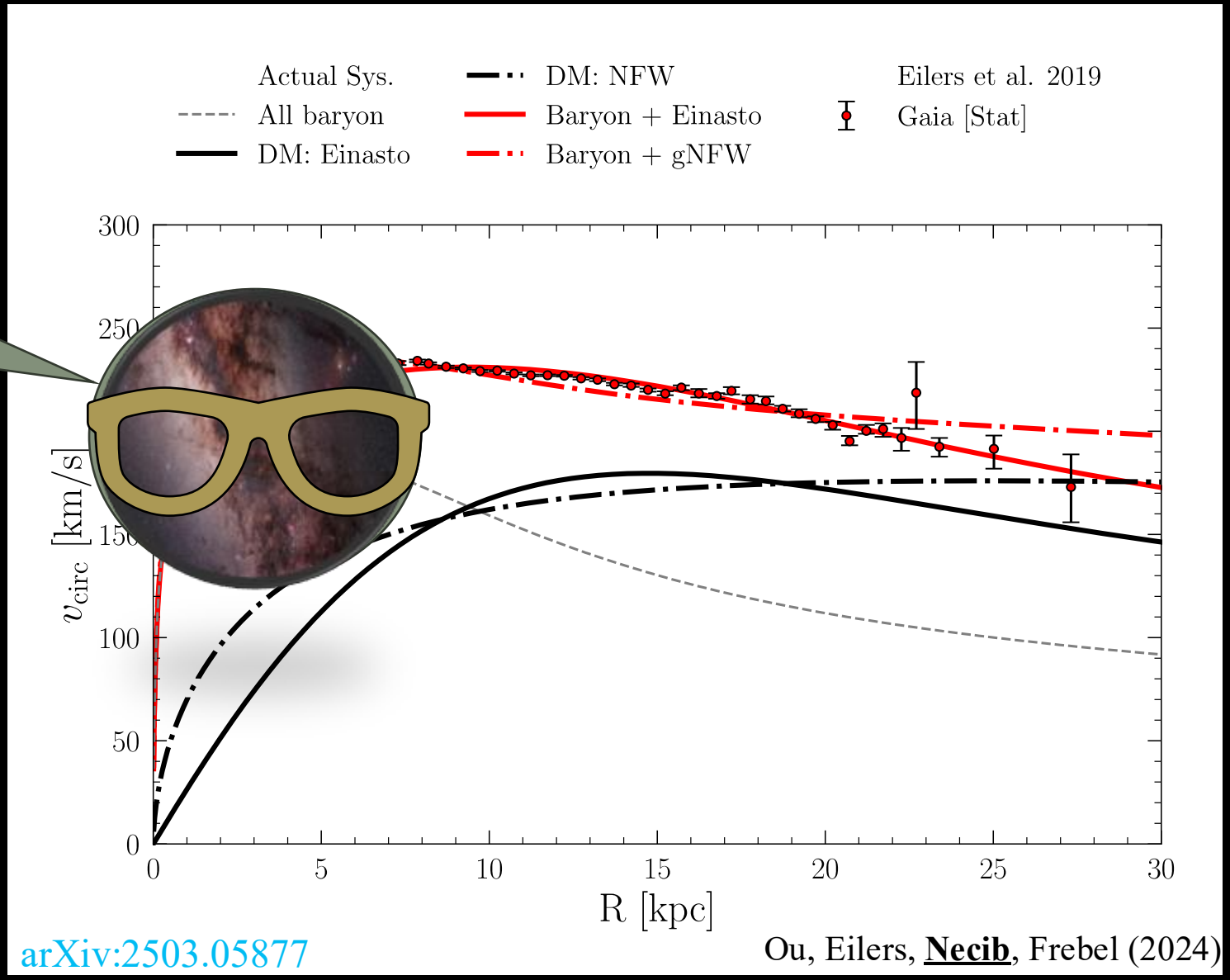
arXiv:2503.05877

Ou, Eilers, **Necib**, Frebel (2024)

Ou, **Necib**, et al. (2025)

WHAT I NEED THEN IS SEE-THROUGH-DUST GLASSES

What I really want, is data here!



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Ou, Eilers, **Necib**, Frebel (2024)

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WE ACTUALLY HAVE SEE-THROUGH-DUST GLASSES...

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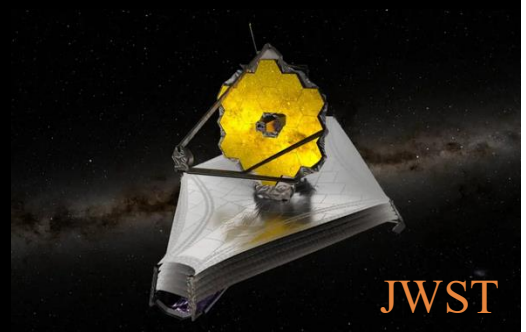

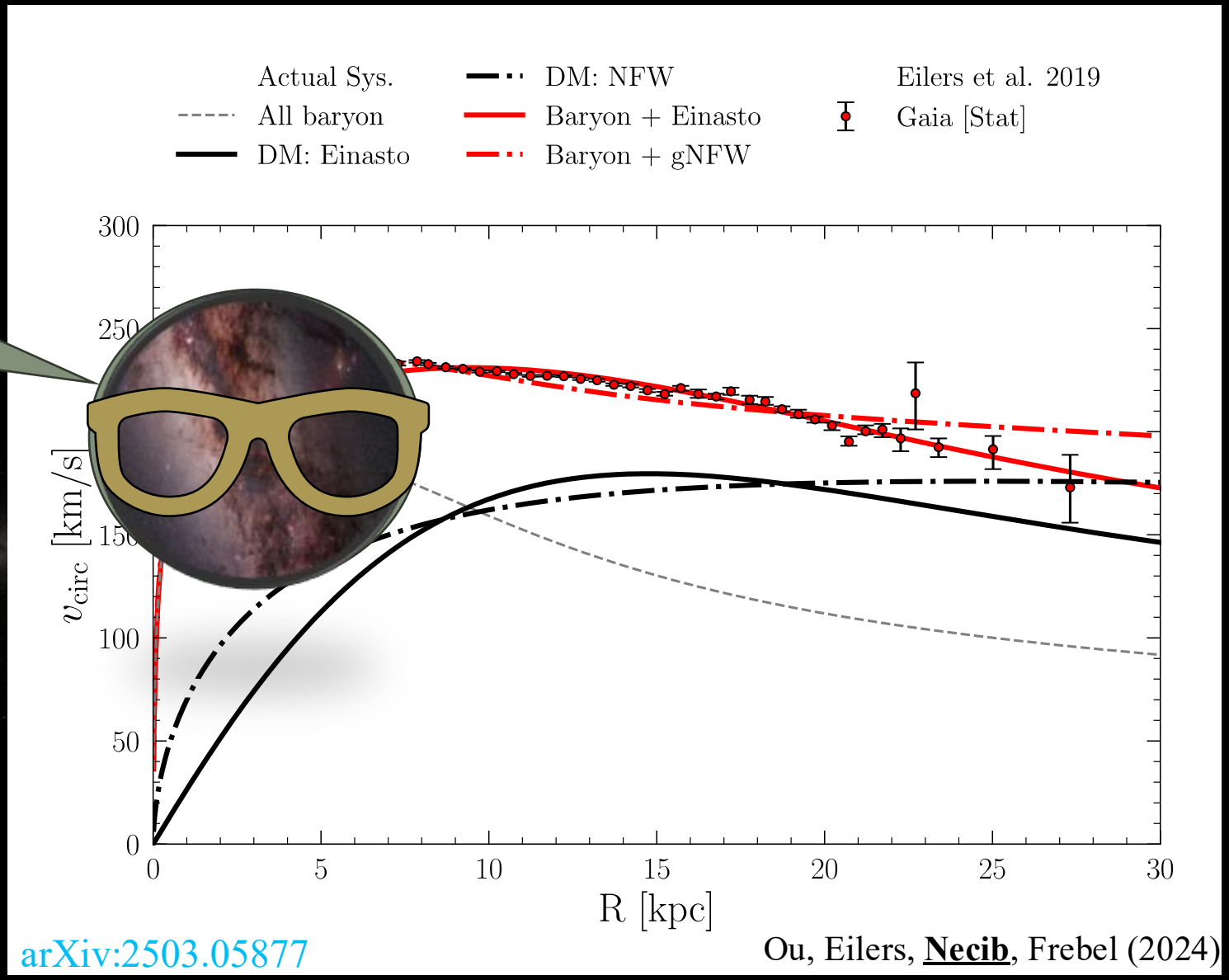

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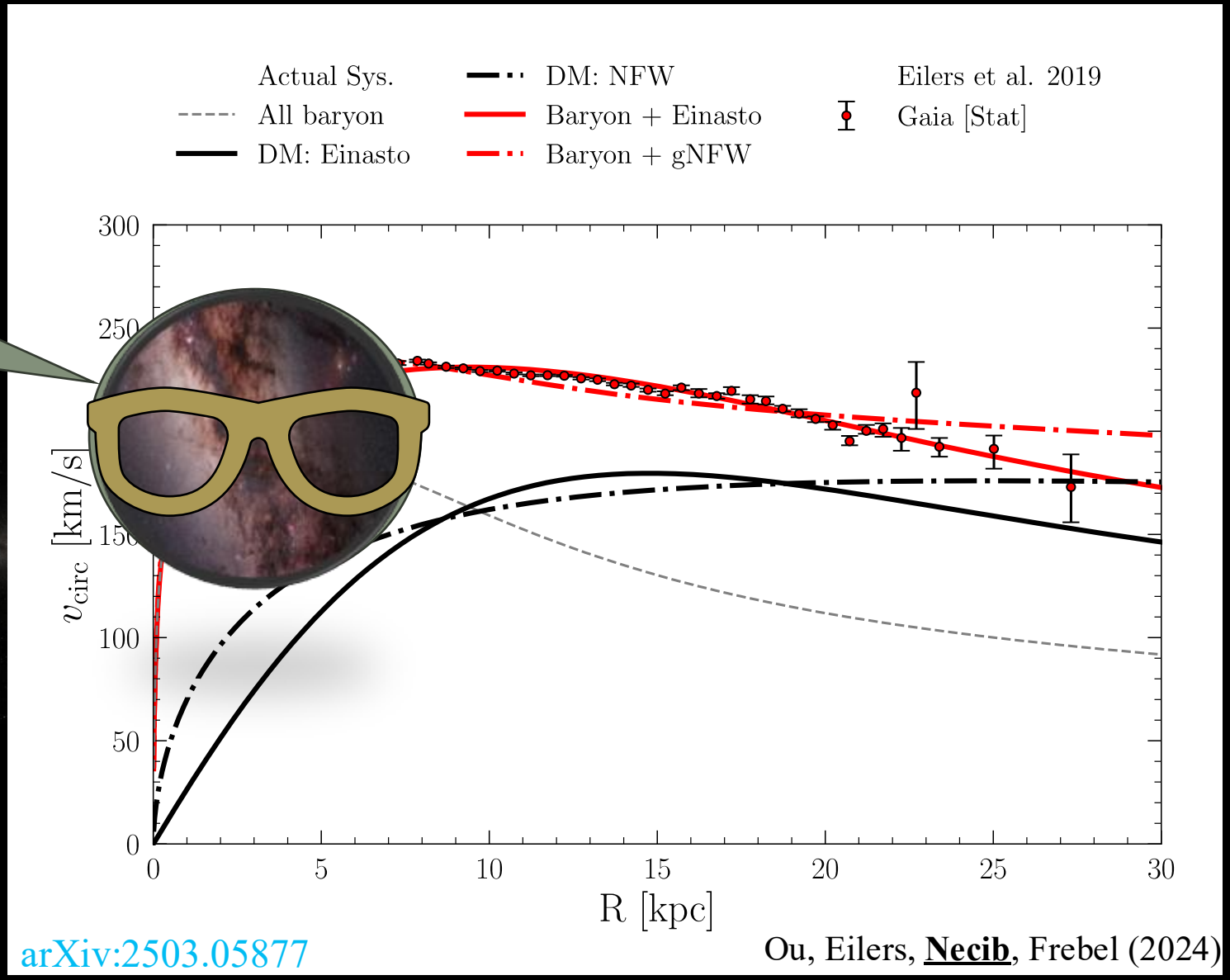

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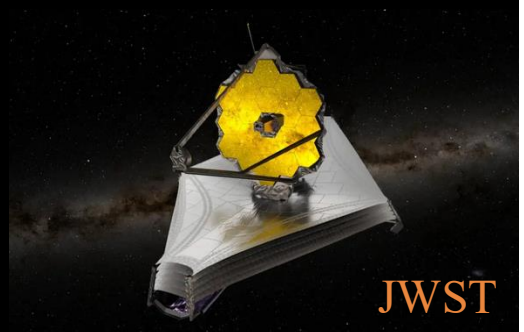


JWST

Roman

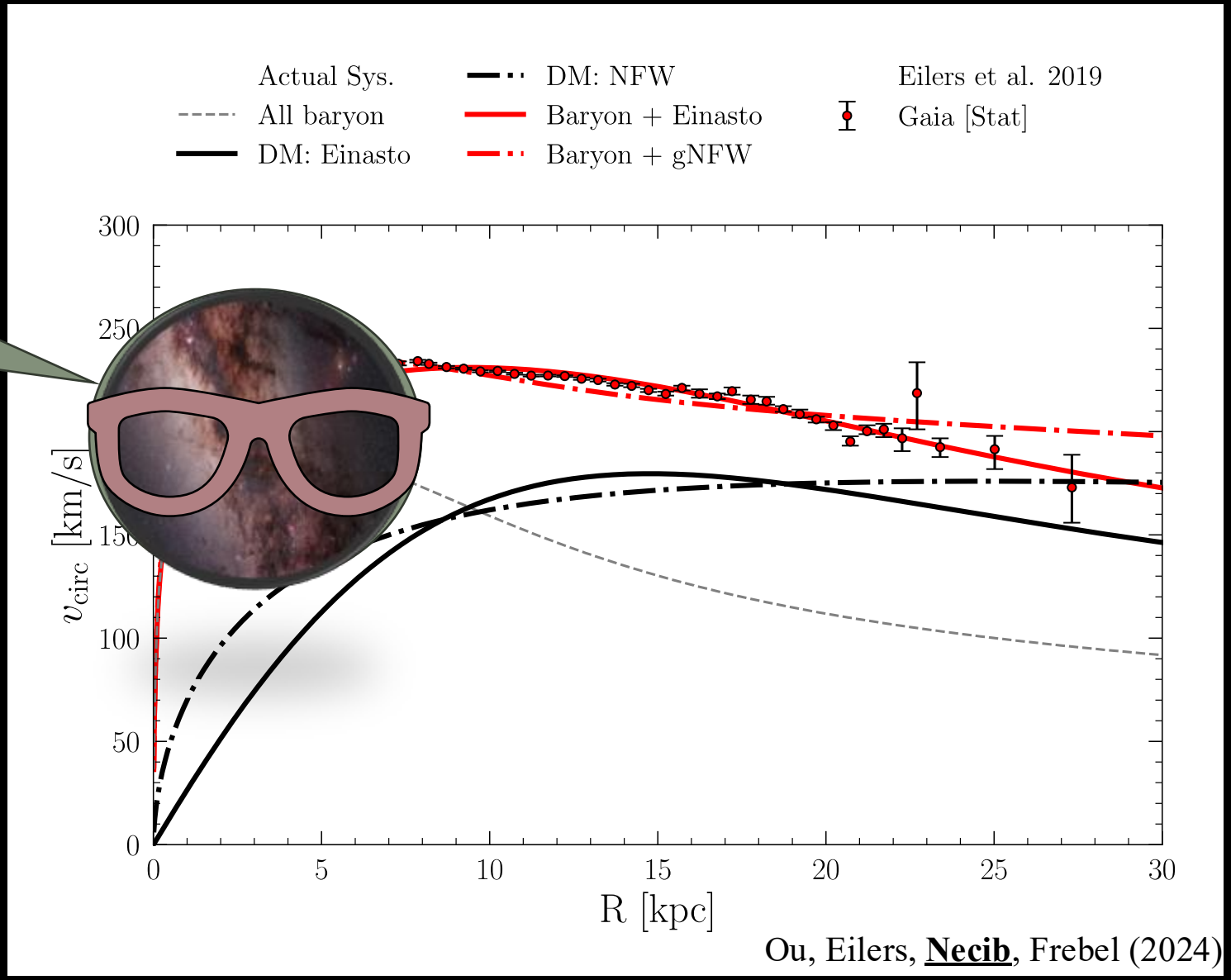
VLT

arXiv:2503.05877

Ou, Eilers, **Necib**, Frebel (2024)

Ou, **Necib**, et al. (2025)

IN THE MEANTIME, LET'S HAVE THEORY GLASSES



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IN THE MEANTIME, LET'S HAVE THEORY GLASSES

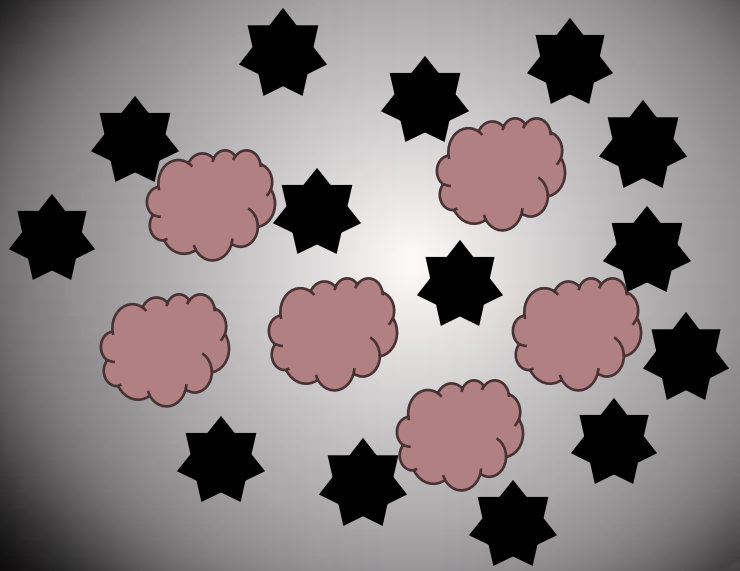
Abdelaziz Hussein

ΛCDM

Lina Necib, MIT

IN THE MEANTIME, LET'S HAVE THEORY GLASSES

Abdelaziz Hussein



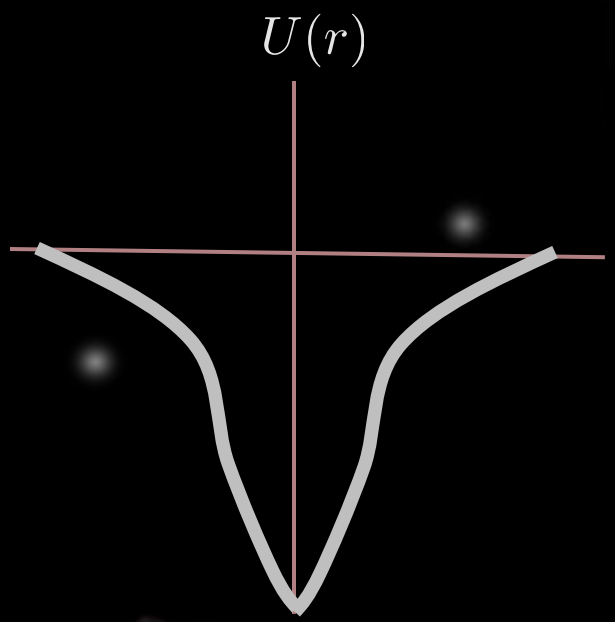
Adiabatic Contraction

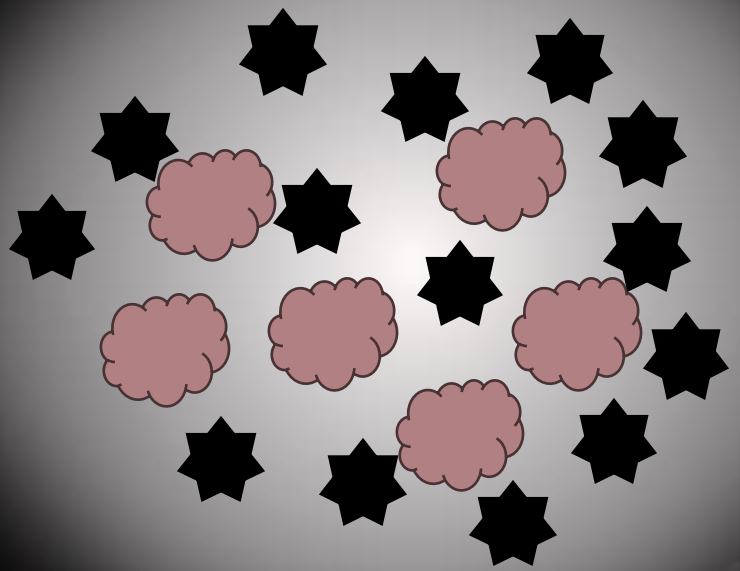
Lina Necib, MIT

Blumenthal et al. 1986
Gnedin et al. 2004

IN THE MEANTIME, LET'S HAVE THEORY GLASSES

Abdelaziz Hussein





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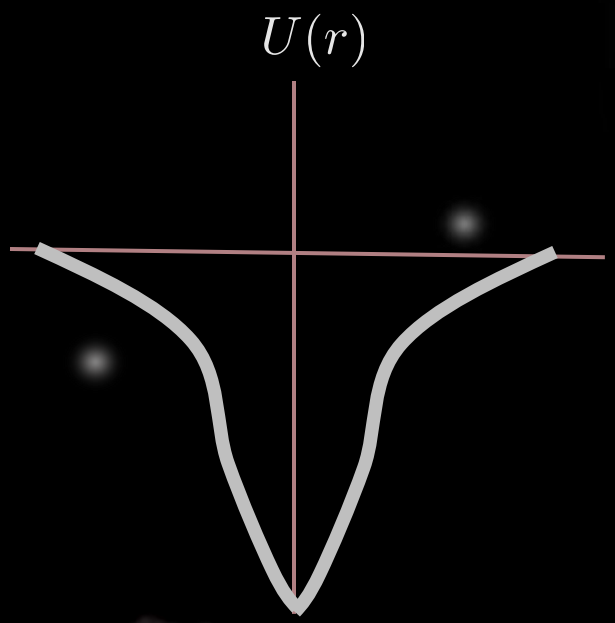

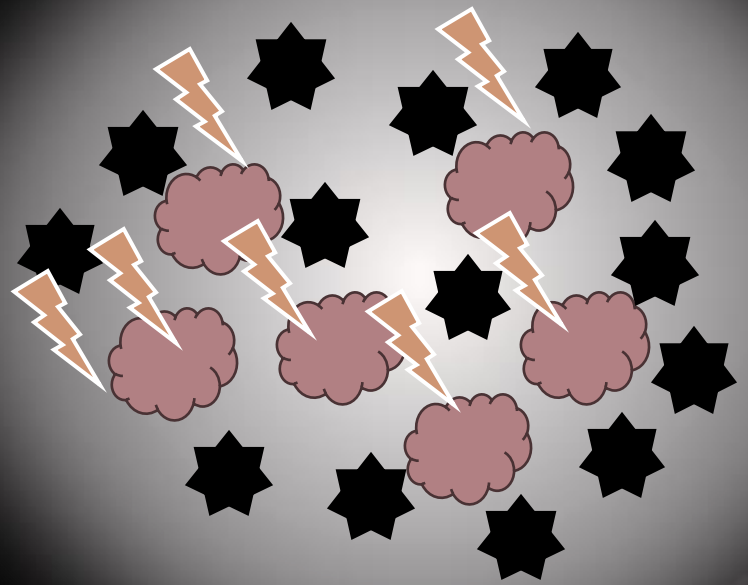

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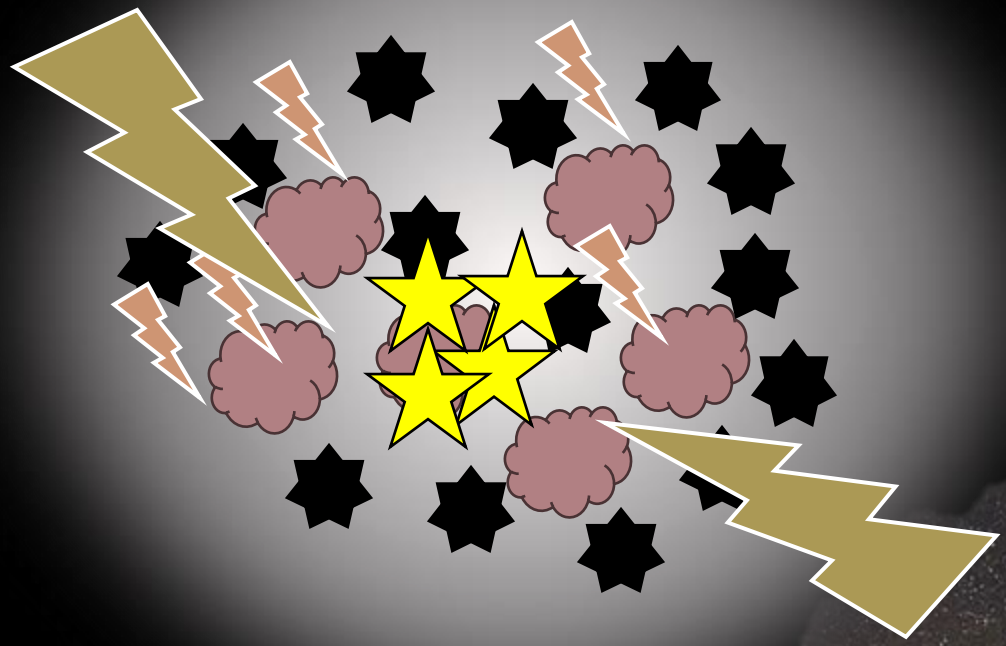

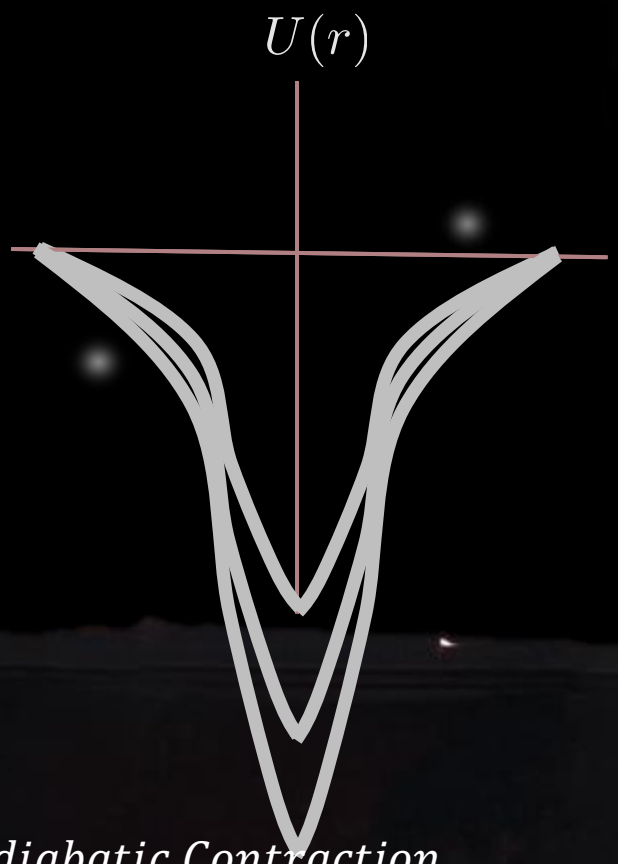

Adiabatic Contraction

Lina Necib, MIT

Blumenthal et al. 1986
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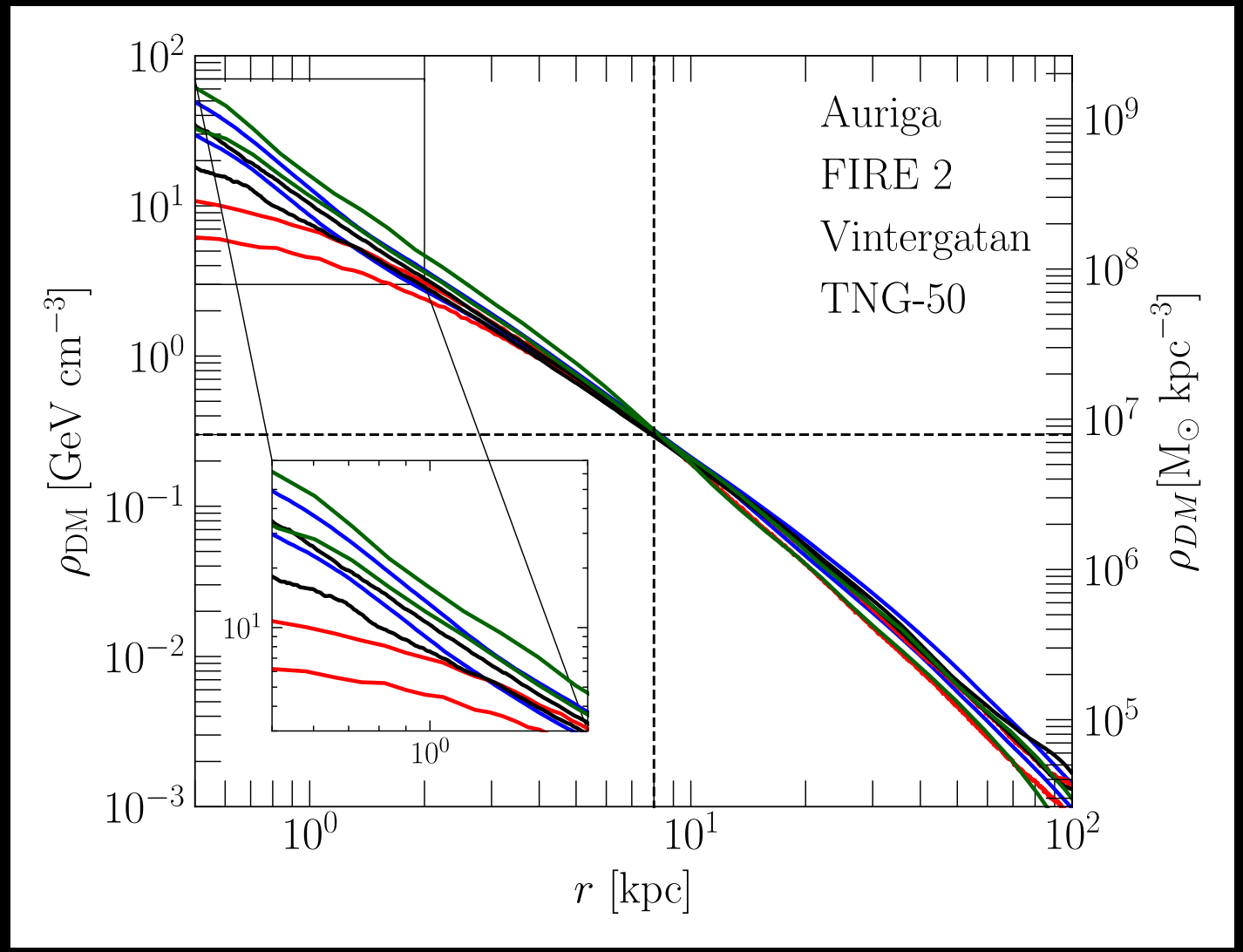
DIFFERENT SIMULATIONS IMPLEMENT PROCESSES AFFECTING THE DENSITY PROFILE DIFFERENTLY

Abdelaziz Hussein



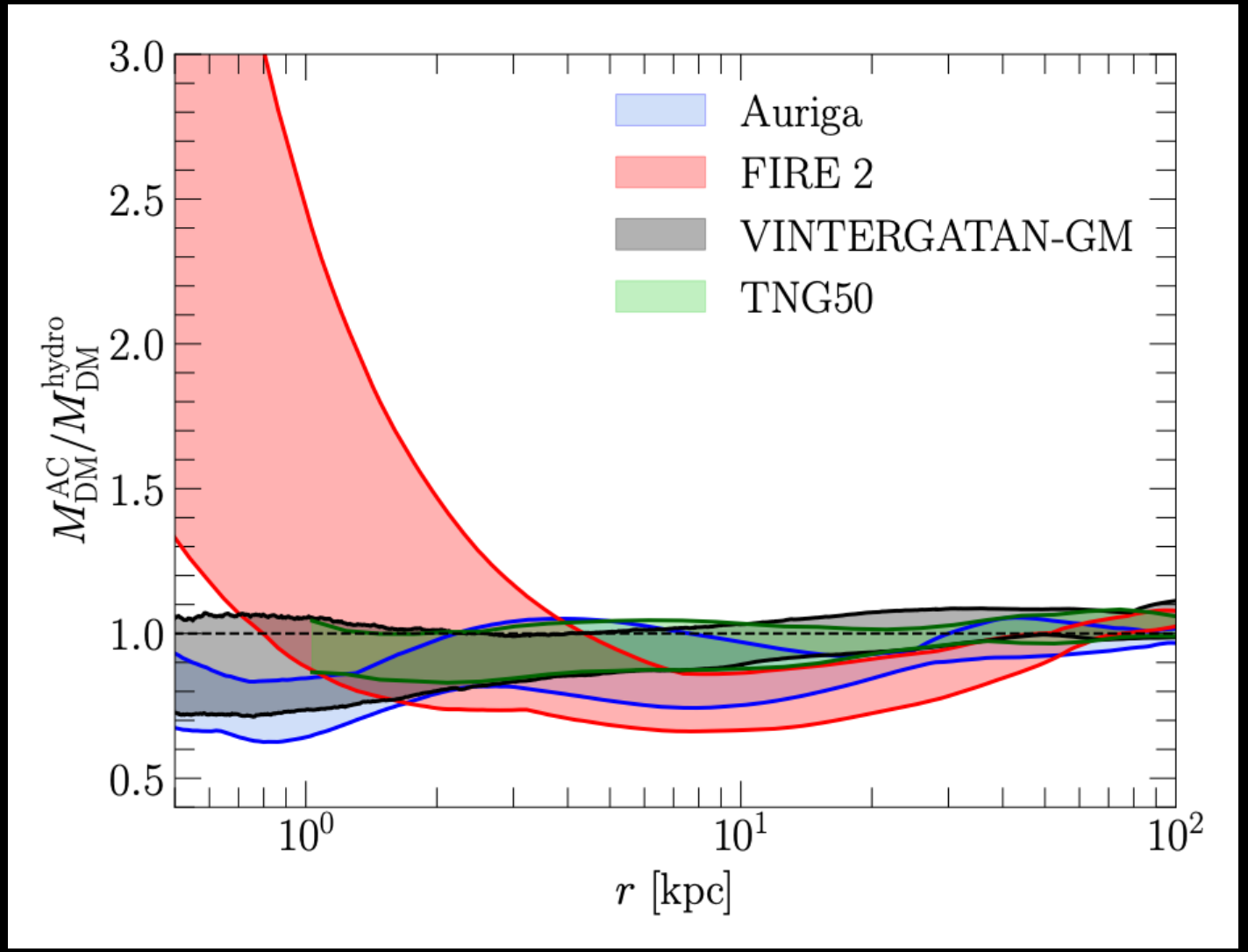
Wetzel et al. (2016)
Hopkins et al. (2018)
Wetzel et al. (2023)
Grand et al. (2024)
Agertz et al. (2021)
Rey et al. (2023)
Pillepich et al. (2023)

Hussein, **Necib**, Kaplinghat et al. (2025) arXiv:2501.14868

Lina Necib, MIT

COMPARING THESE RESULTS TO THE THEORY PREDICTION LEADS TO

Abdelaziz Hussein



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Lina Necib, MIT

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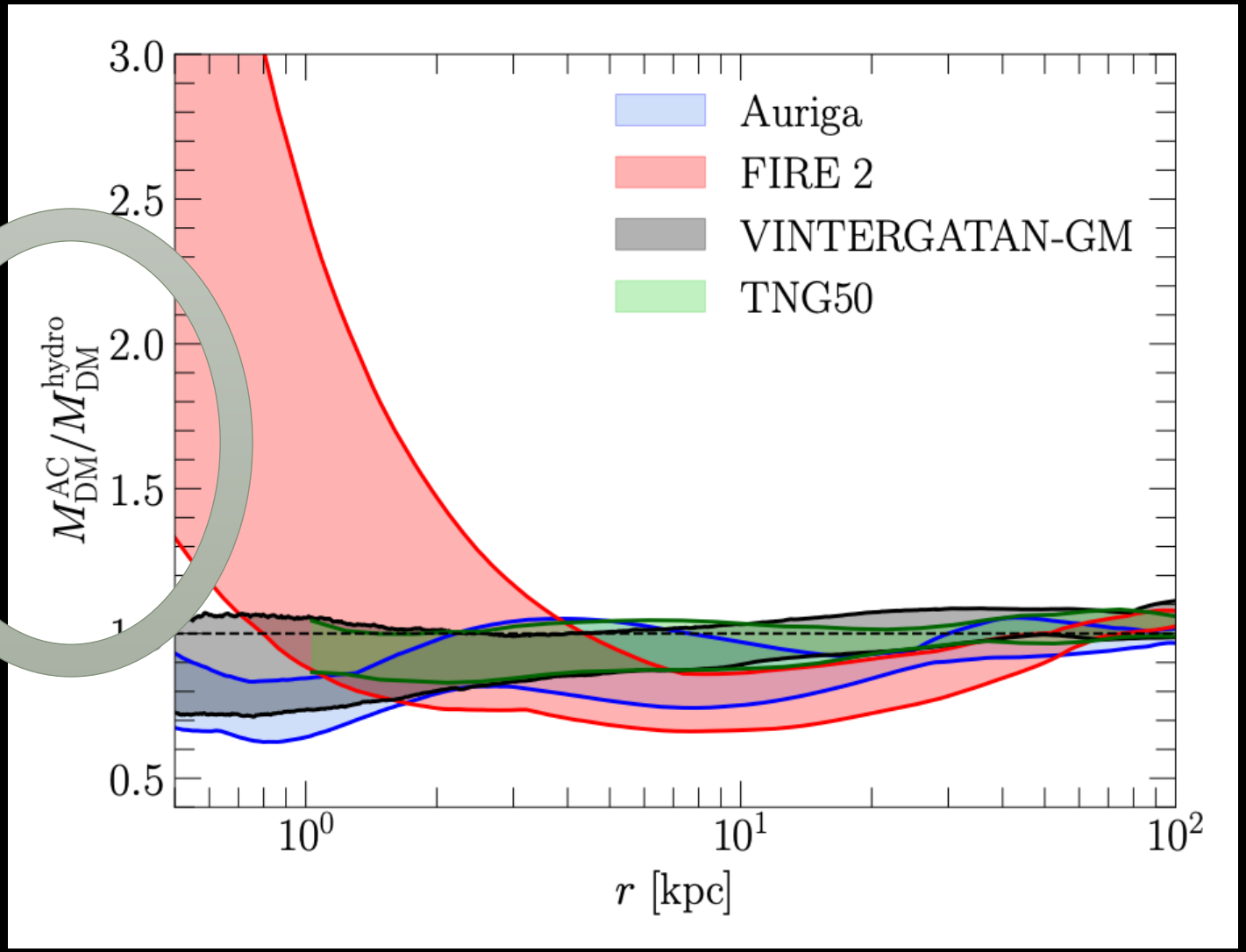
Ratio of theory expectation to simulation results



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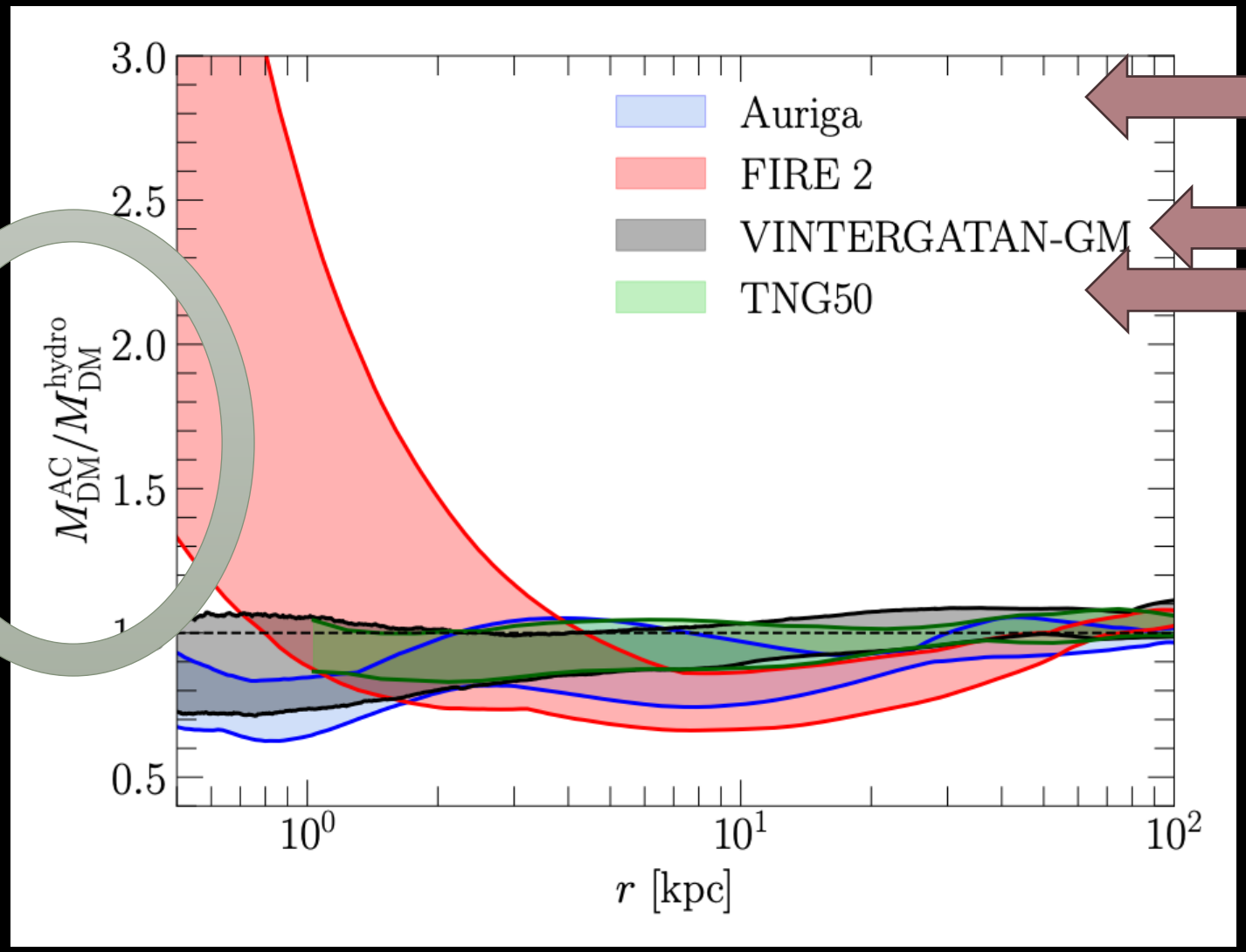

Adiabatic Contraction is a good description of the model

Wetzel et al. (2016)
Hopkins et al. (2018)
Wetzel et al. (2023)
Grand et al. (2024)
Agertz et al. (2021)
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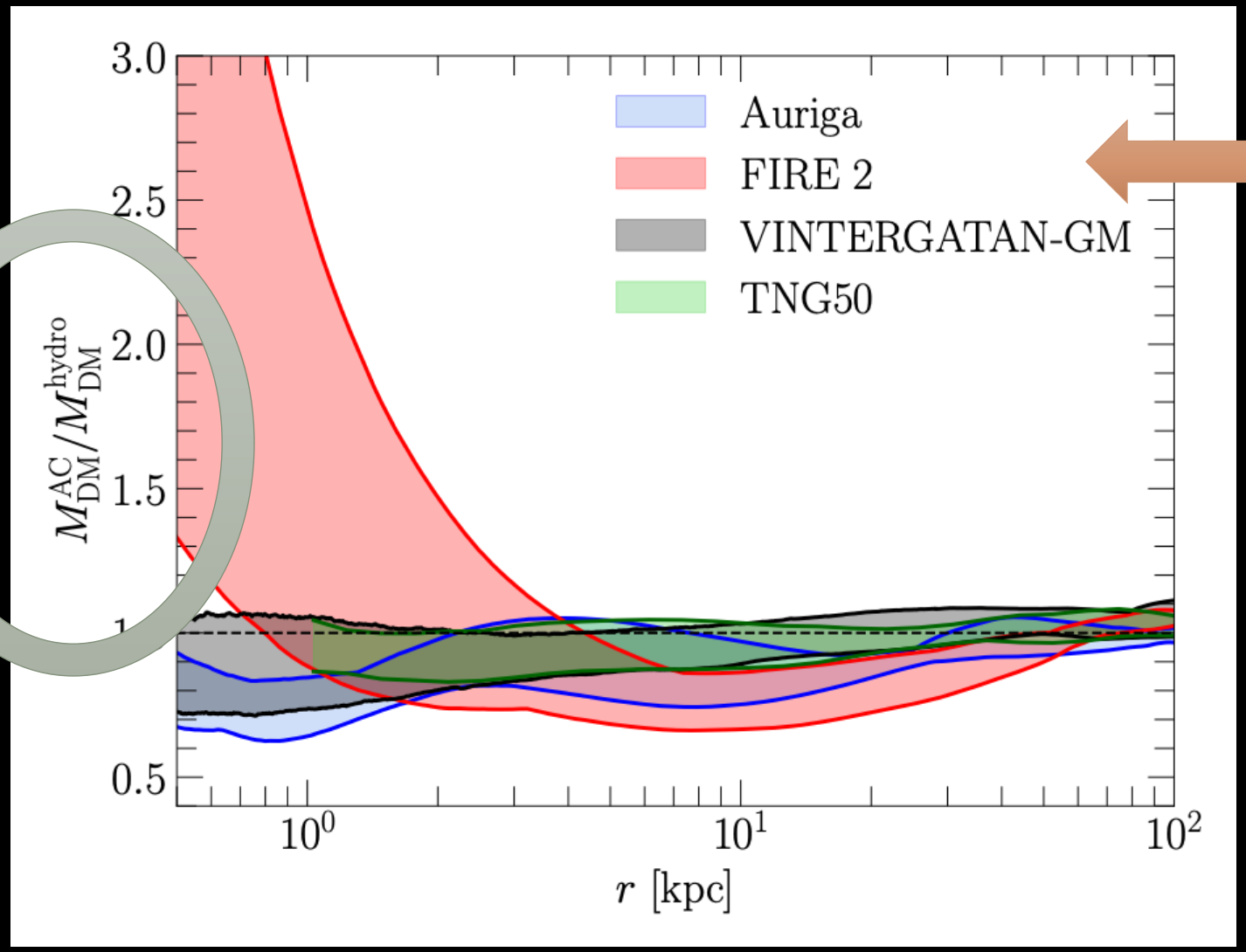
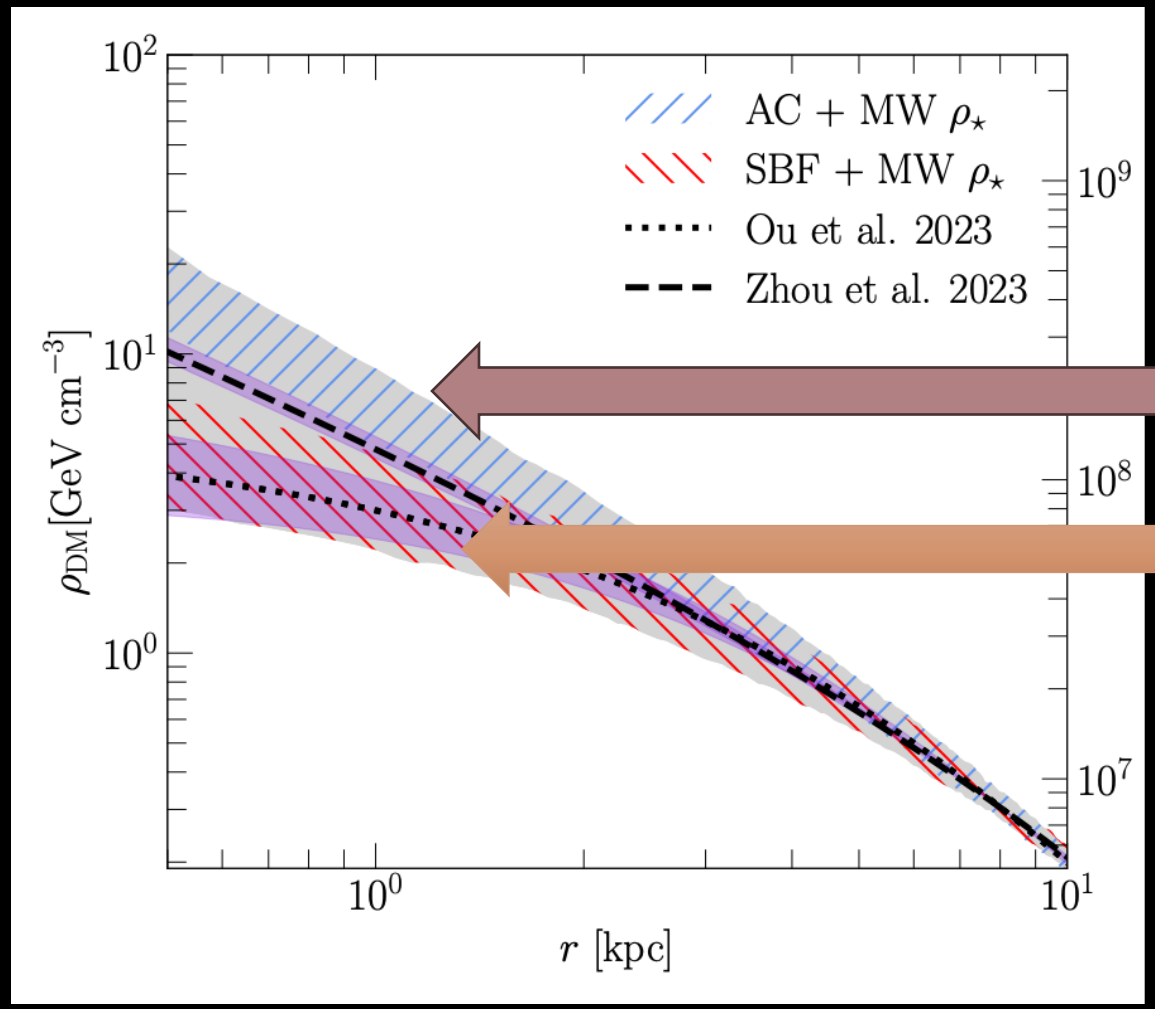


Strong Feedback is breaking the adiabatic contraction model

Wetzel et al. (2016)
Hopkins et al. (2018)
Wetzel et al. (2023)
Grand et al. (2024)
Agertz et al. (2021)
Rey et al. (2023)
Pillepich et al. (2023)

LET'S QUANTIFY OUR IGNORANCE BY SETTING A SYSTEMATIC BAND FOR THE MILKY WAY

Abdelaziz Hussein



Adiabatic Contraction density profile based on the Milky Way observed stellar distribution

Strong Baryonic Feedback model from FIRE-2, assuming the Milky Way observed stellar distribution

Hussein, **Necib**, Kaplinghat et al. (2025) arXiv:2501.14868

Lina Necib, MIT

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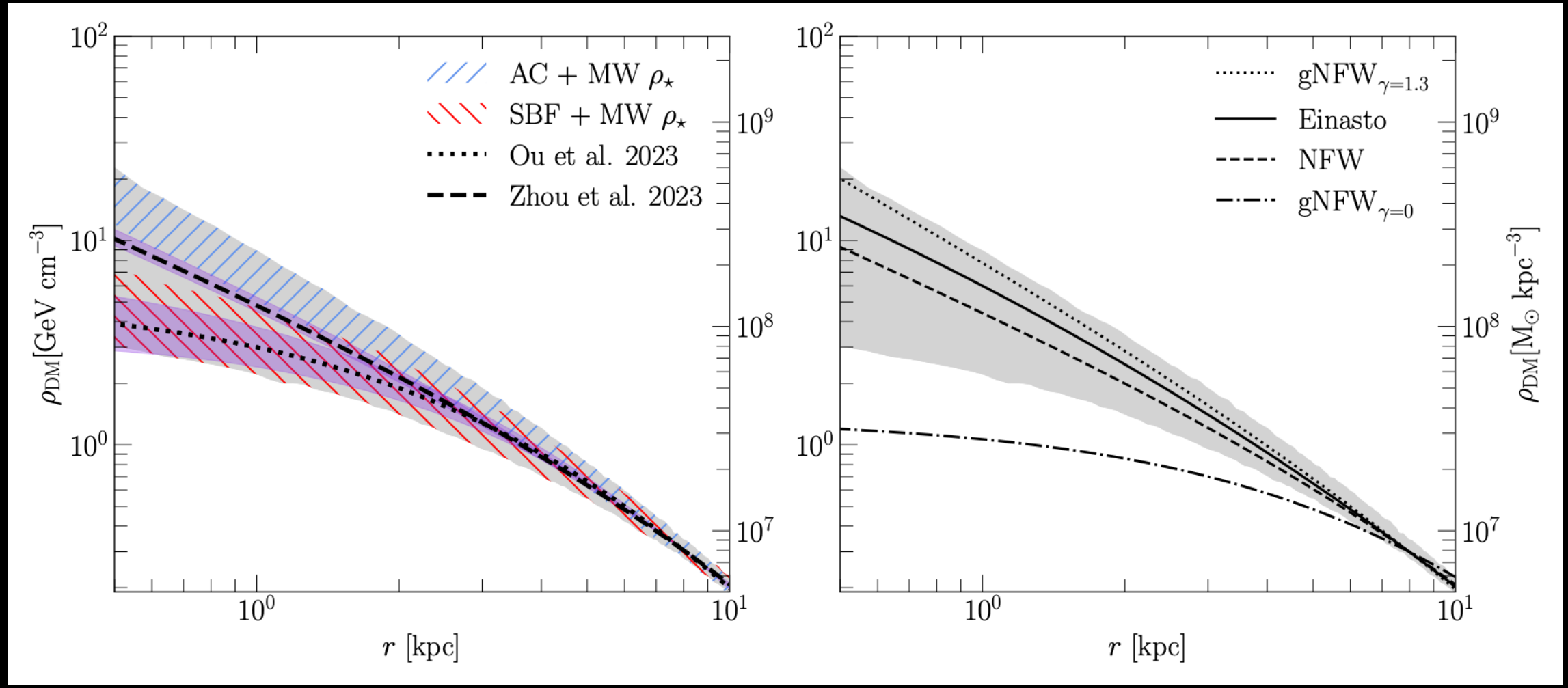

Hussein, **Necib**, Kaplinghat et al. (2025) arXiv:2501.14868

Lina Necib, MIT

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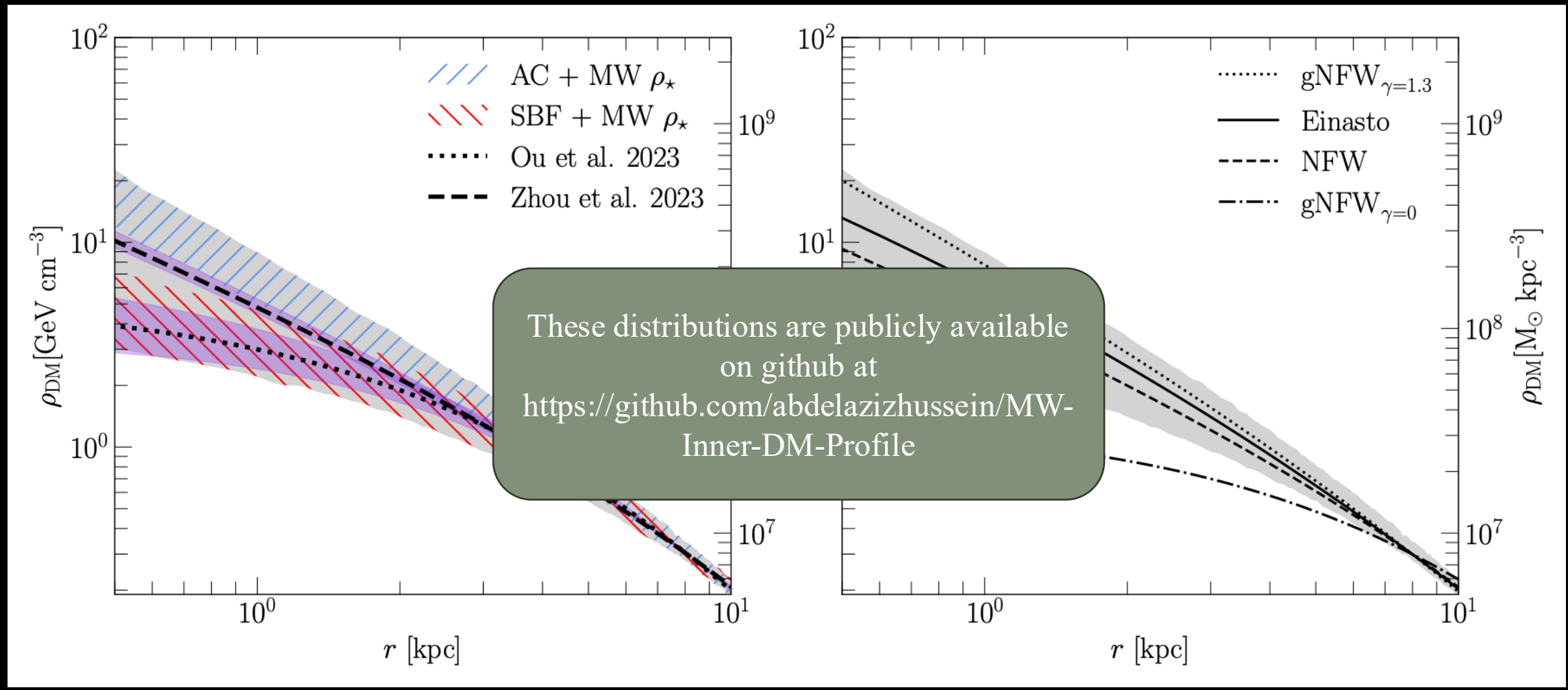


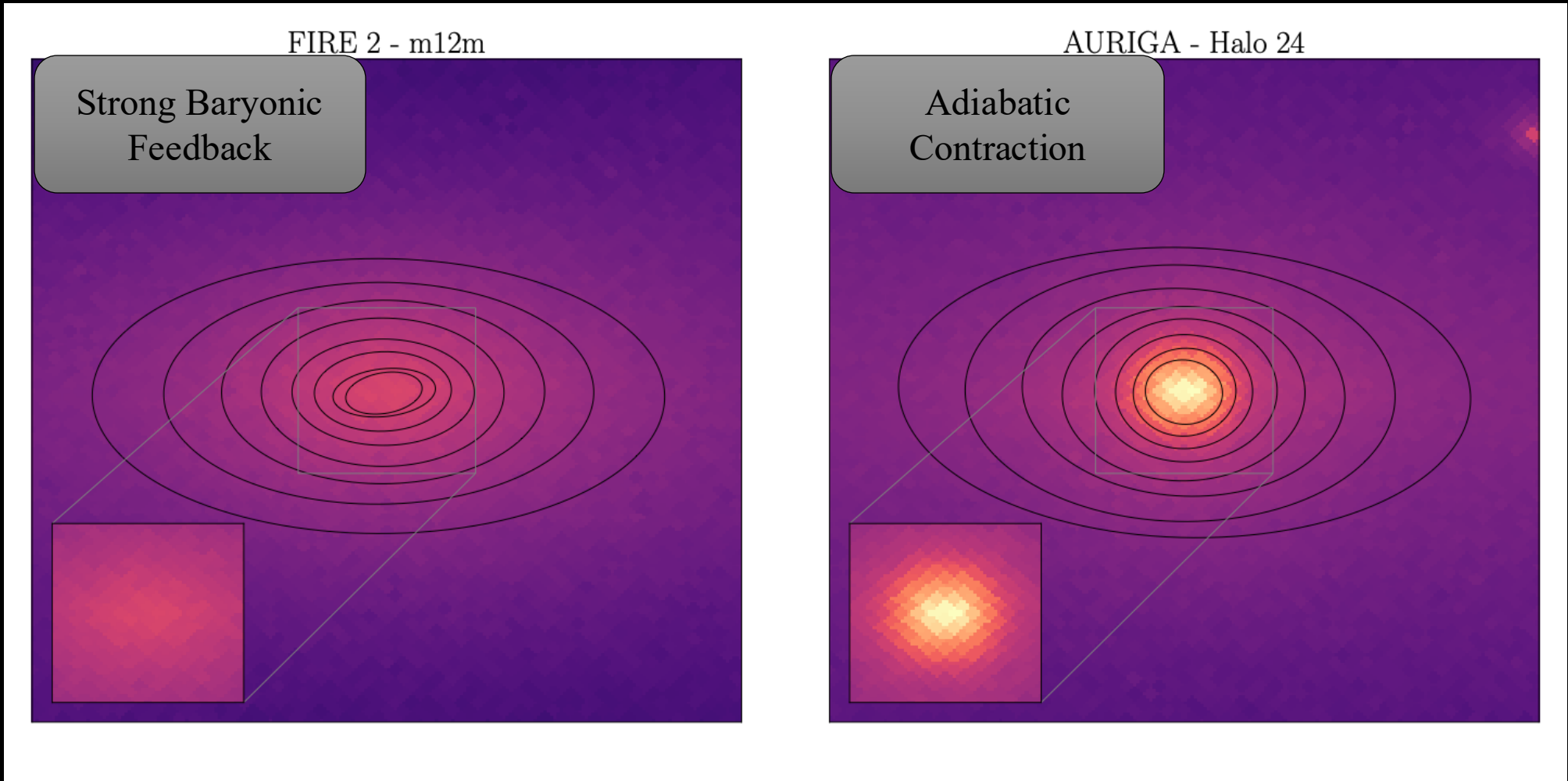
These distributions are publicly available on github at https://github.com/abdelazizhussein/MW-Inner-DM-Profile

Hussein, **Necib**, Kaplinghat et al. (2025) arXiv:2501.14868

Lina Necib, MIT

THE DIFFERENT MECHANISMS ALSO AFFECT THE DARK MATTER SPATIAL DISTRIBUTION

Abdelaziz Hussein

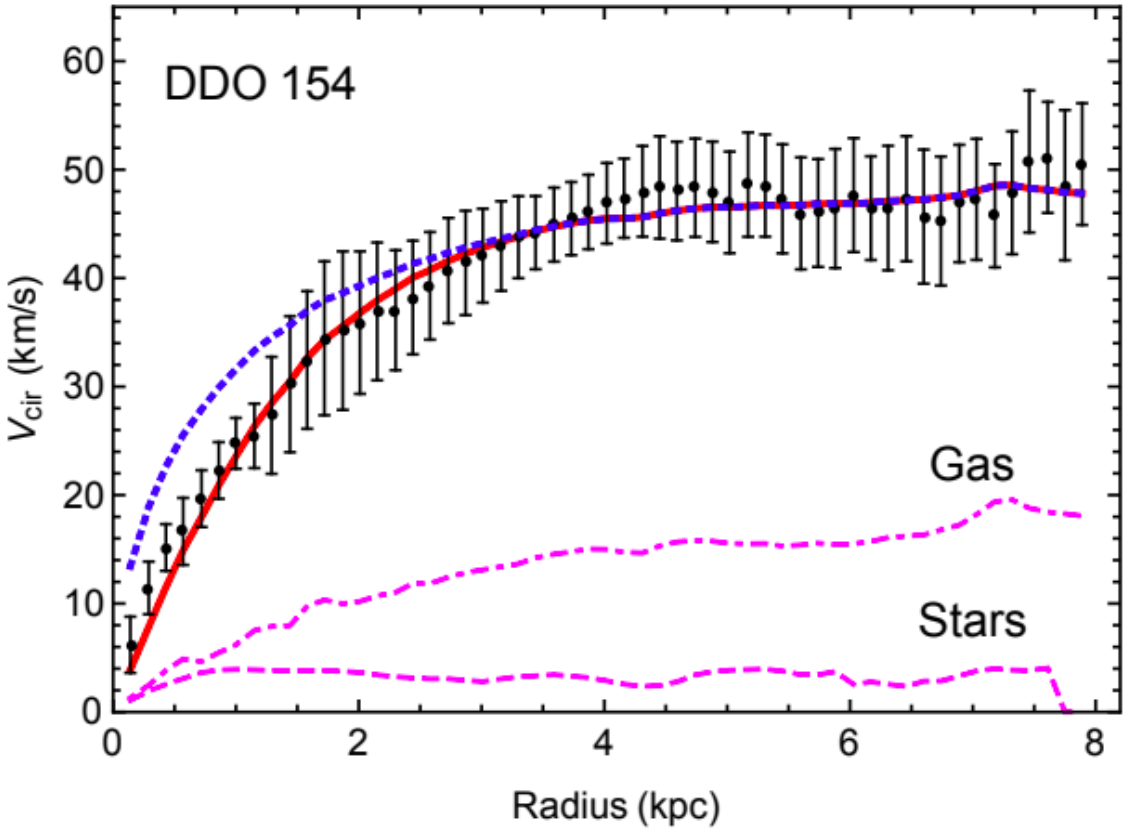


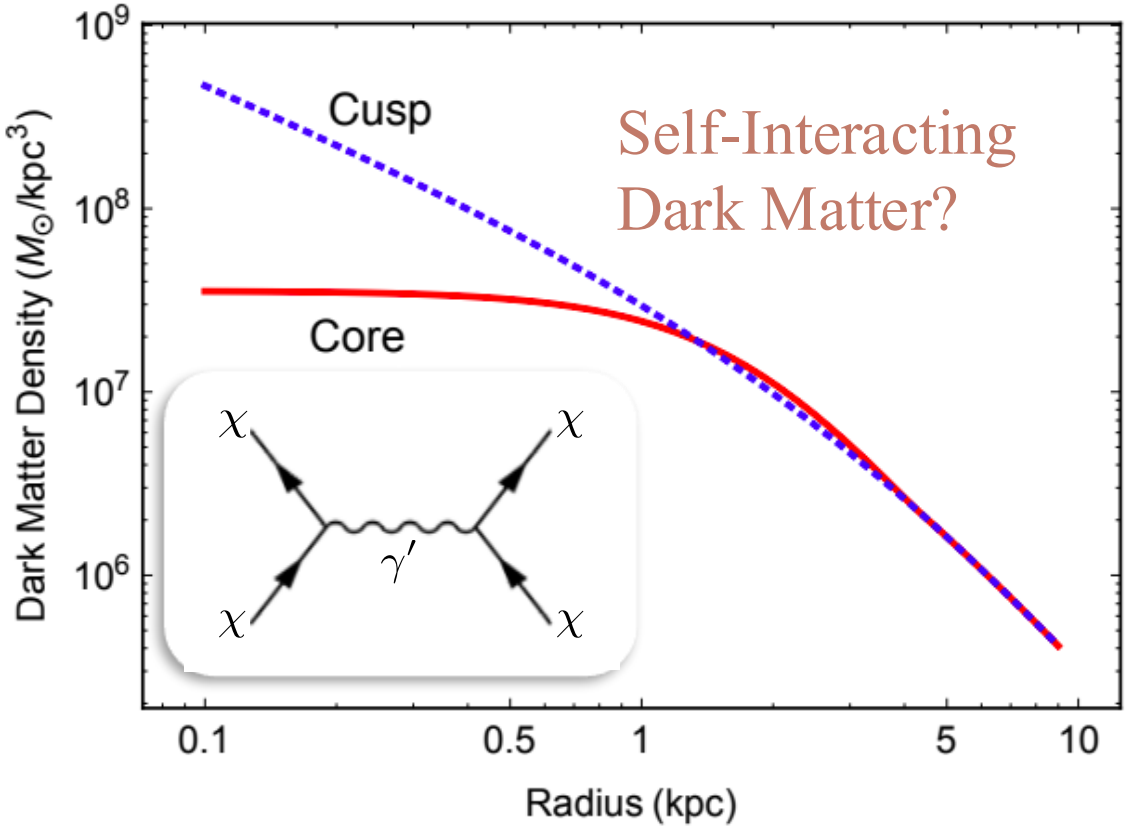
Hussein, **Necib**, Kaplinghat et al. (2025) arXiv:2501.14868

Lina Necib, MIT

TODAY'S TALK:
MAPPING THE DARK MATTER DENSITY PROFILE IN THE MILKY WAY AND DWARF GALAXIES

EARLY WORKS HAVE SHOWN THAT DWARF GALAXIES MIGHT HAVE CORES





Lina Necib, MIT

Spergel & Steinhardt (2000)

Oh et al. (2015)
Tulin & Yu (2017)

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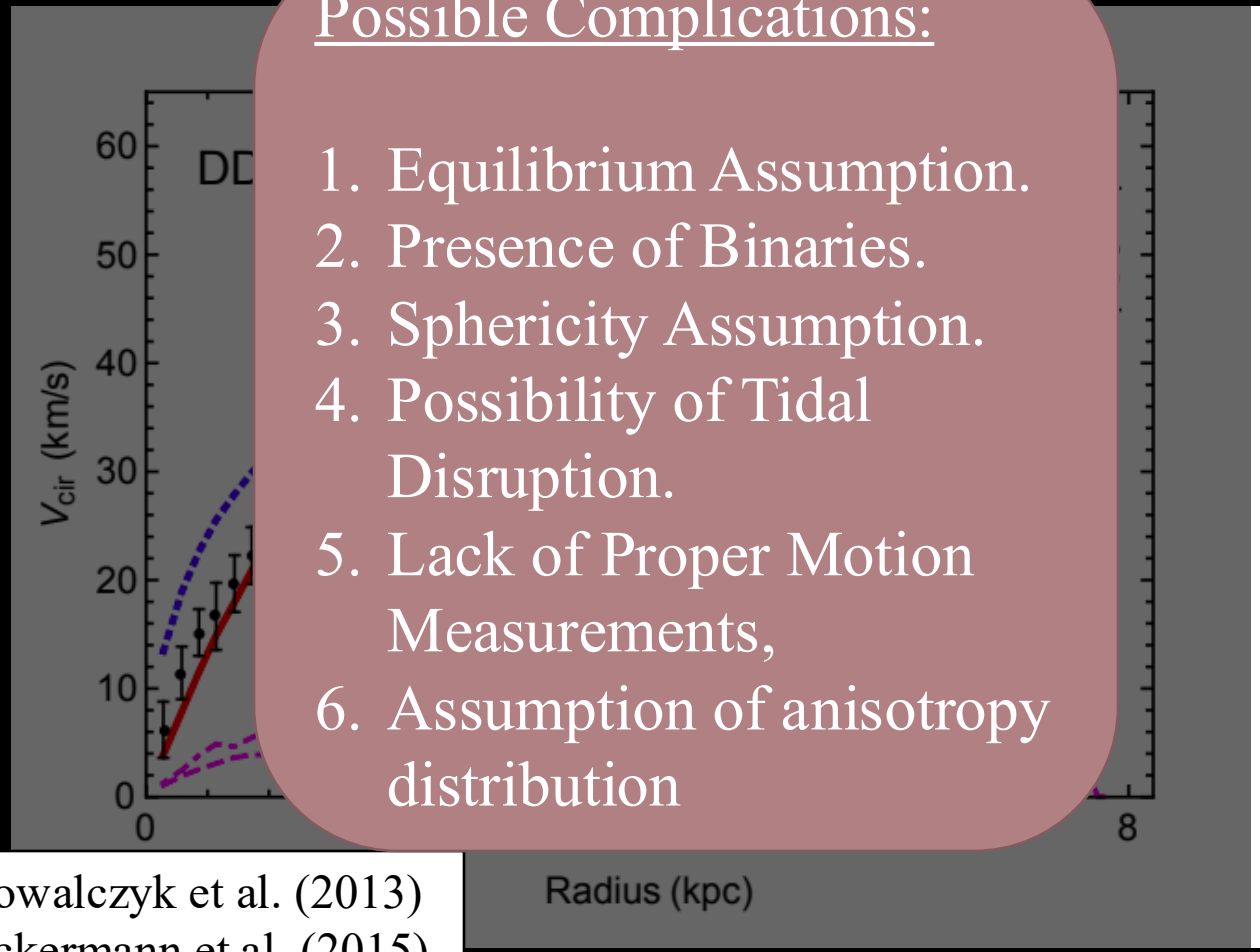
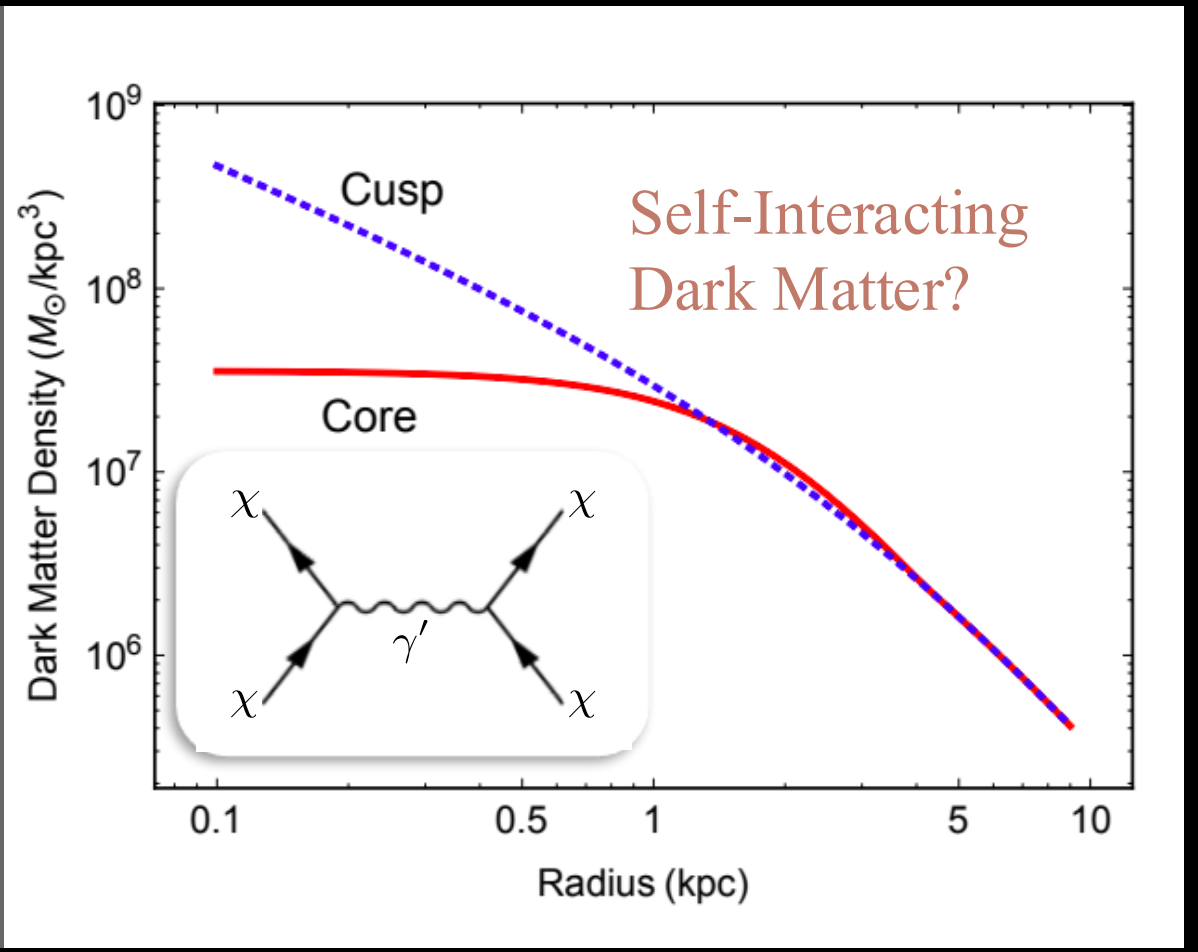


Possible Complications:

1. Equilibrium Assumption.
2. Presence of Binaries.
3. Sphericity Assumption.
4. Possibility of Tidal Disruption.
5. Lack of Proper Motion Measurements,
6. Assumption of anisotropy distribution



Kowalczyk et al. (2013)
Ackermann et al. (2015)
Bonnivard et al. (2015)
El-Badry et al. (2017)

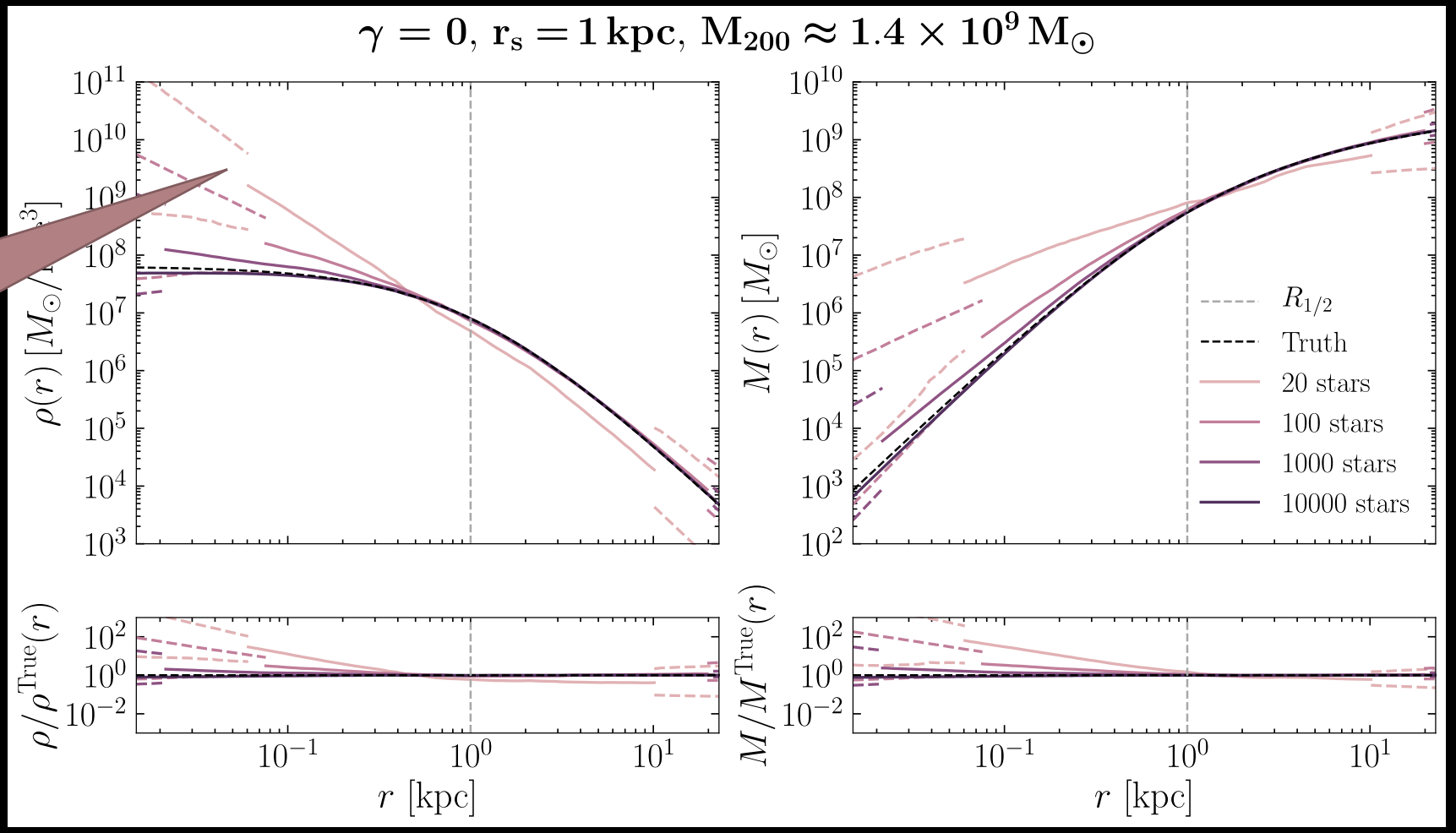
Lina Necib, MIT

Spergel & Steinhardt (2000)

Oh et al. (2015)
Tulin & Yu (2017)

ONE NEEDS ABOUT 10,000 STARS TO GET THE PROFILE CORRECTLY IN THE STANDARD JEANS ANALYSIS

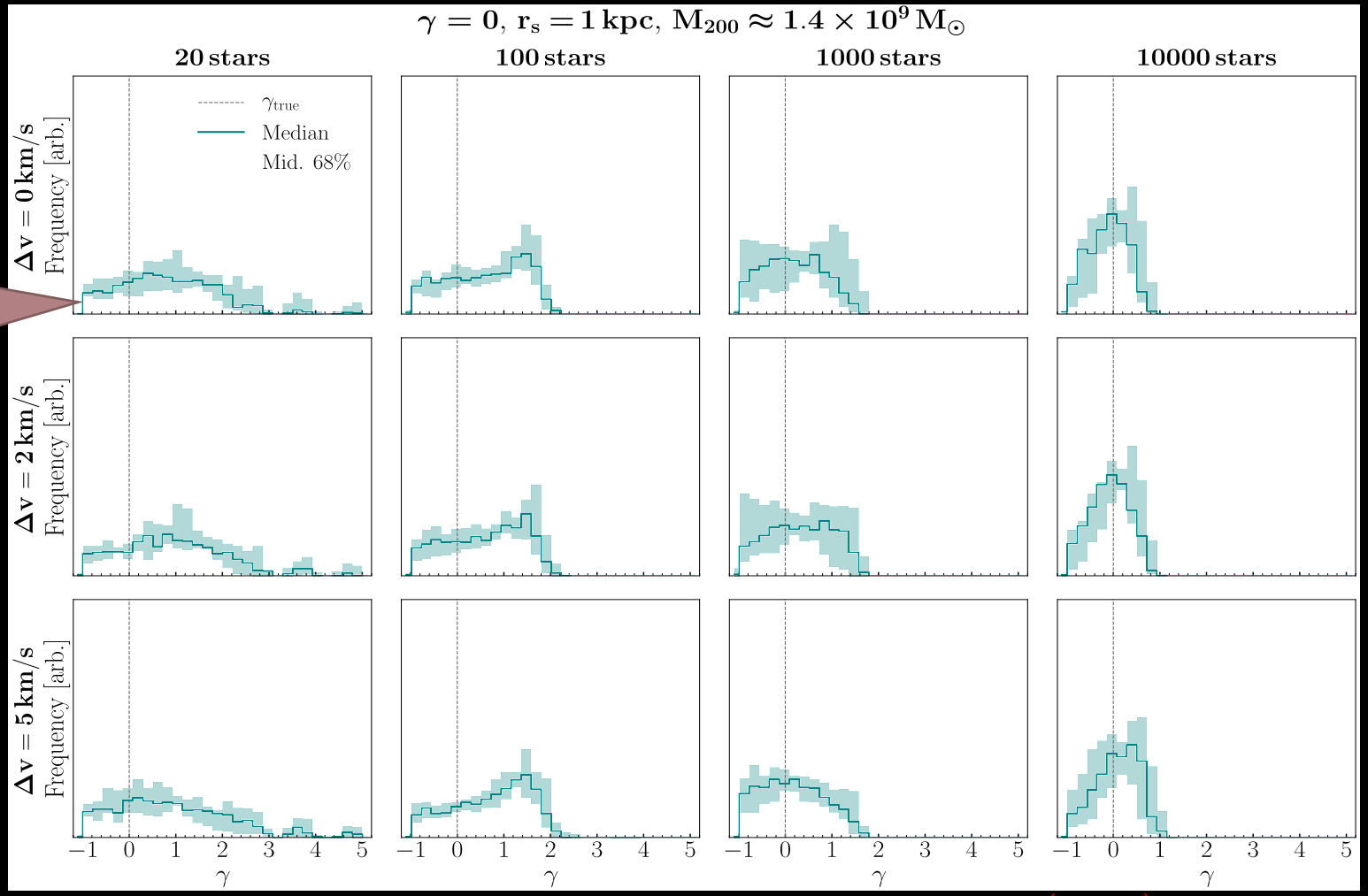
Cores are harder to resolve with fewer statistics



Lina Necib, MIT

Chang & **Necib** (2020)

ONE NEEDS ABOUT 10,000 STARS TO GET THE PROFILE CORRECTLY IN THE STANDARD JEANS ANALYSIS



Cores are harder to resolve with fewer statistics

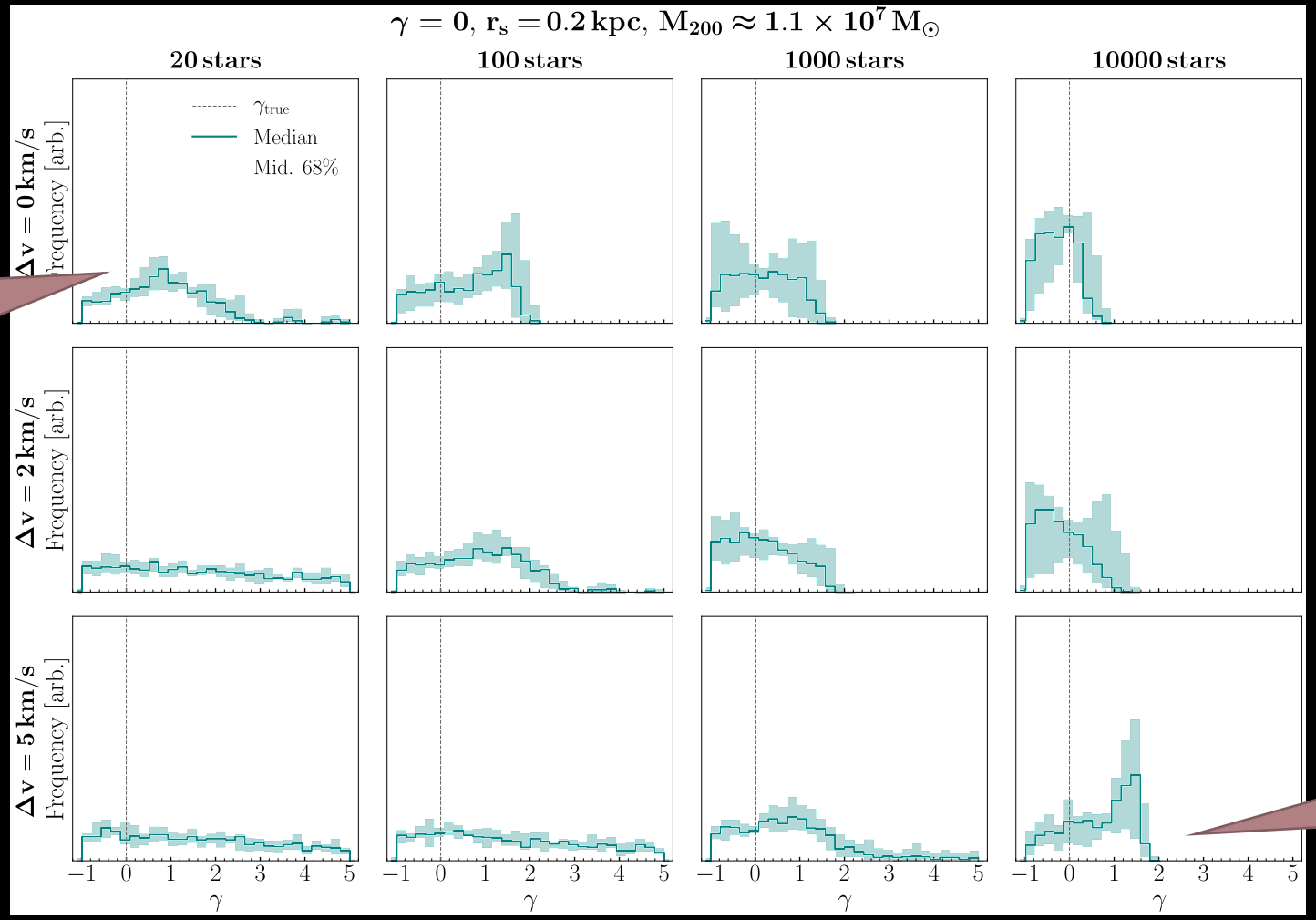
Cusp: $\gamma = 1$
Core: $\gamma = 0$

$M_{200} \approx 10^9 M_\odot$

$$\rho = \rho_0 \left(\frac{r}{r_s}\right)^{-\gamma} \left(1 + \frac{r}{r_s}\right)^{-(3-\gamma)}$$

Chang & **Necib** (2020)

ONE NEEDS ABOUT 10,000 STARS TO GET THE PROFILE CORRECTLY IN THE STANDARD JEANS ANALYSIS



Cores are harder to resolve with fewer statistics

Cusp: $\gamma = 1$
Core: $\gamma = 0$

The errors are worse for the ultrafaints

$M_{200} \approx 10^7 M_\odot$

$$\rho = \rho_0 \left(\frac{r}{r_s}\right)^{-\gamma} \left(1 + \frac{r}{r_s}\right)^{-(3-\gamma)}$$

Chang & **Necib** (2020)

Goal:

Build a method where we can get the correct inner density profile of Dark Matter.

Issue:

Solution:

Goal:

Build a method where we can get the correct inner density profile of Dark Matter.

Issue:

Jeans analysis, the standard method, makes a lot of assumptions to deal with this problem, and requires a LOT of statistics to get the answer right.

Solution:

Goal: Build a method where we can get the correct inner density profile of Dark Matter.

Issue: Jeans analysis, the standard method, makes a lot of assumptions to deal with this problem, and requires a LOT of statistics to get the answer right.

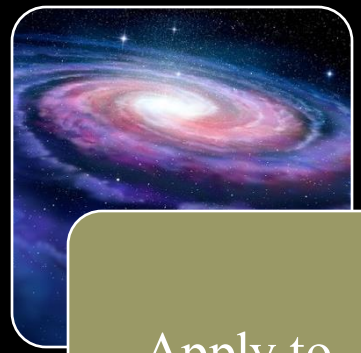
Solution: Use simulation based inference + Graph Neural Networks to train a model that obtains the density profiles of dwarfs.

WE USE SIMULATION-BASED INFERENCE TO OBTAIN THE INNER PROFILE OF THE DARK MATTER FROM THE KINEMATICS OF THE STARS

Generate Idealized Simulations

Build the Graph to establish relationships between the stars

Train the Network



Apply to data

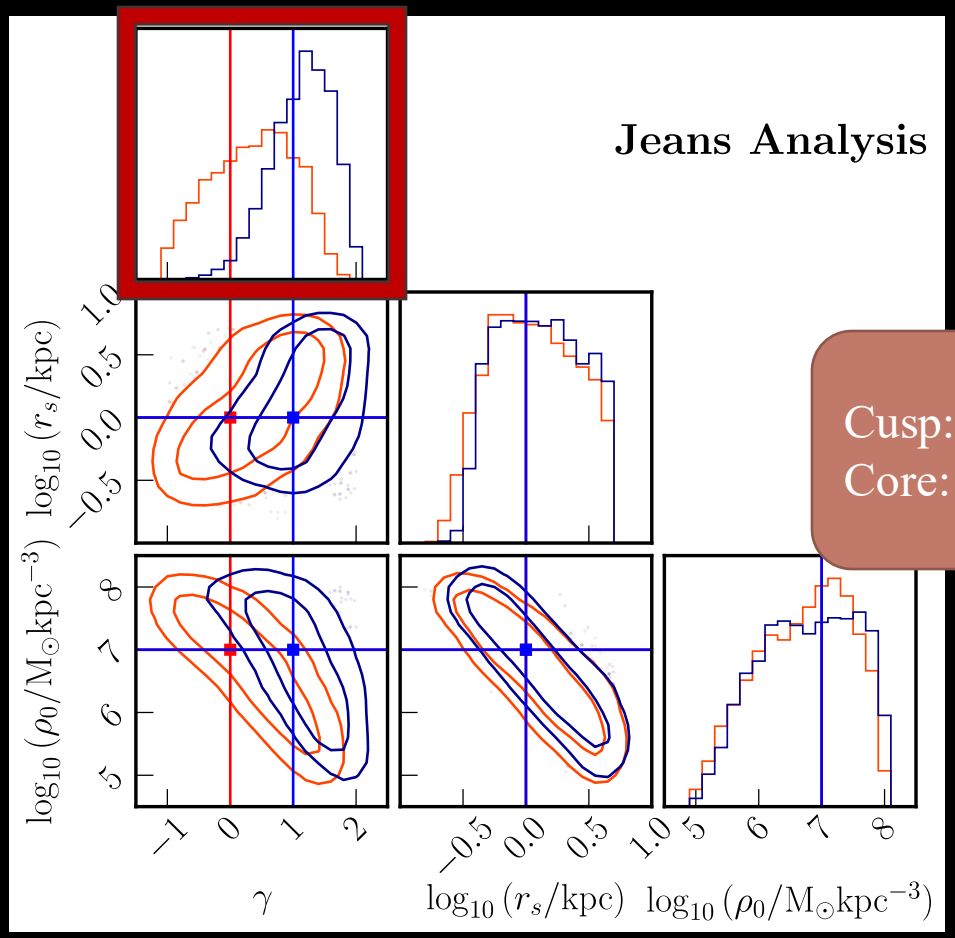
Use Normalizing Flow to obtain a posterior distribution

Nguyen, Mishra-Sharma, Williams, & **Necib** (2022)

GNN + NORMALIZING FLOWS OUTPERFORMS JEANS ANALYSIS

Tri Nguyen

$$\rho = \rho_0 \left(\frac{r}{r_s}\right)^{-\gamma} \left(1 + \frac{r}{r_s}\right)^{-(3-\gamma)}$$



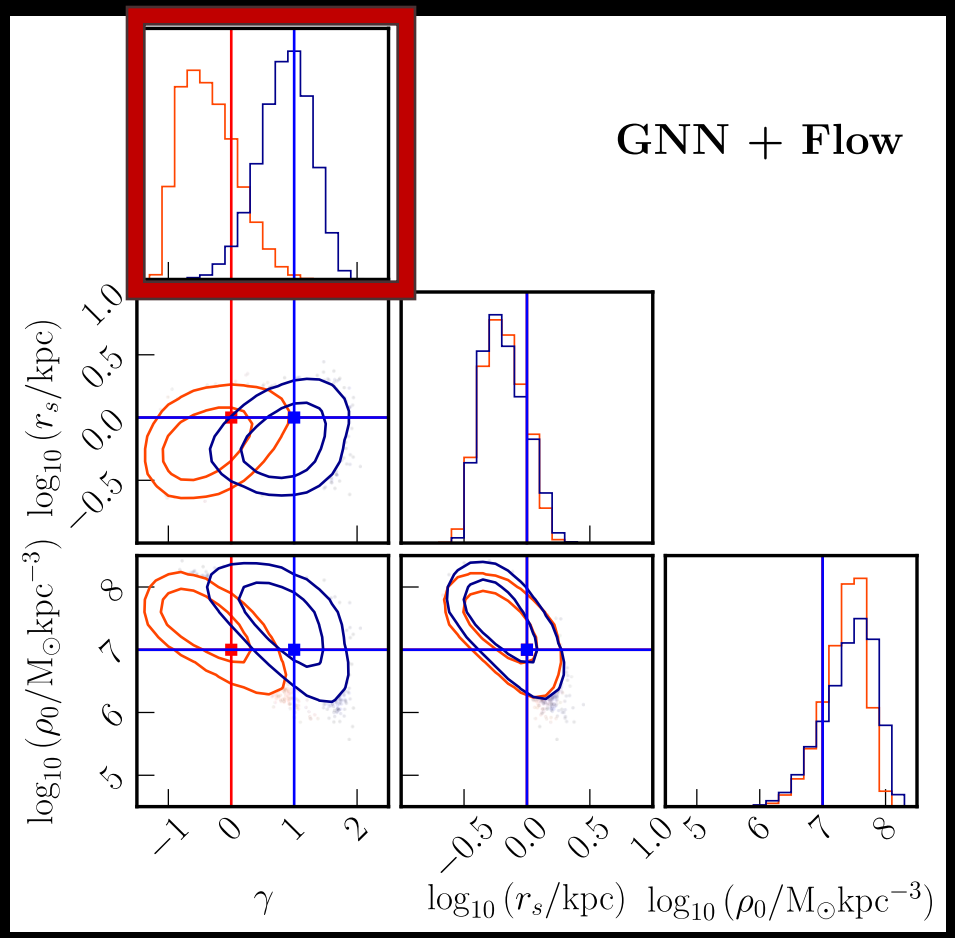
Cusp: $\gamma = 1$
Core: $\gamma = 0$

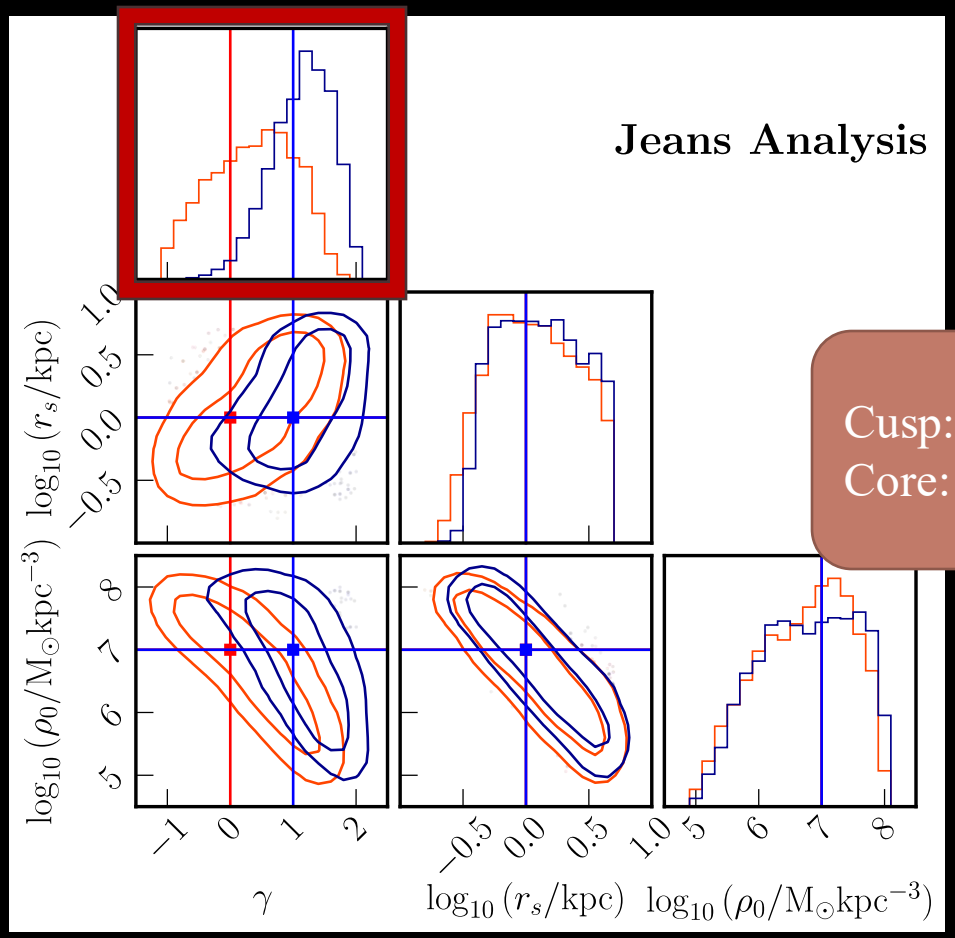
Lina Necib, MIT

Nguyen, Mishra-Sharma, Williams, & **Necib** (2022)

GNN + NORMALIZING FLOWS OUTPERFORMS JEANS ANALYSIS

Tri Nguyen





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Lina Necib, MIT

Nguyen, Mishra-Sharma, Williams, & **Necib** (2022)

GNN + NORMALIZING FLOWS OUTPERFORMS JEANS ANALYSIS

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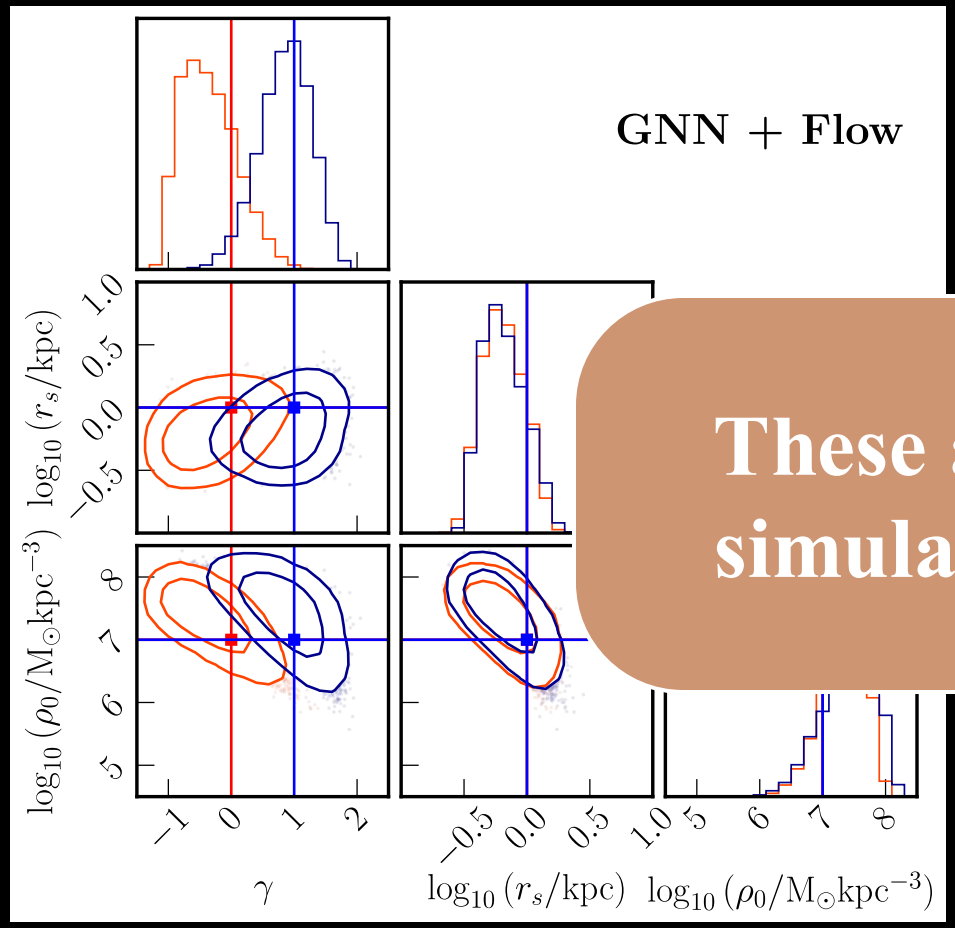
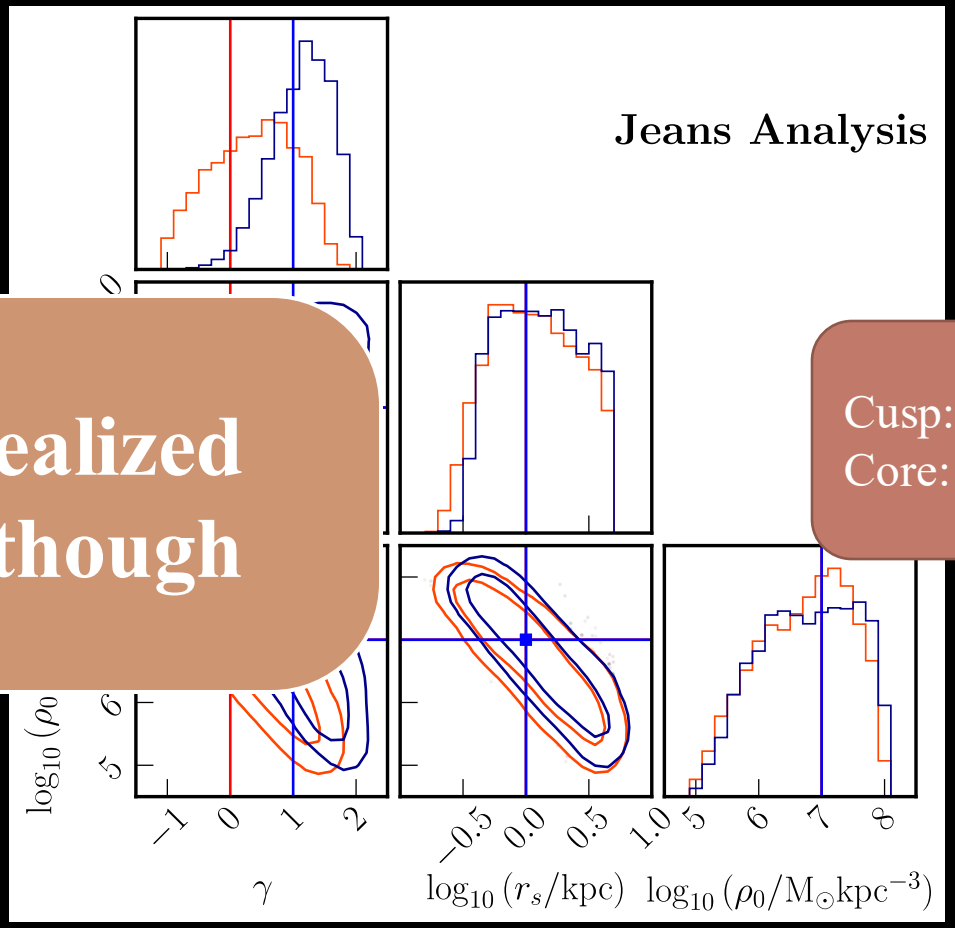


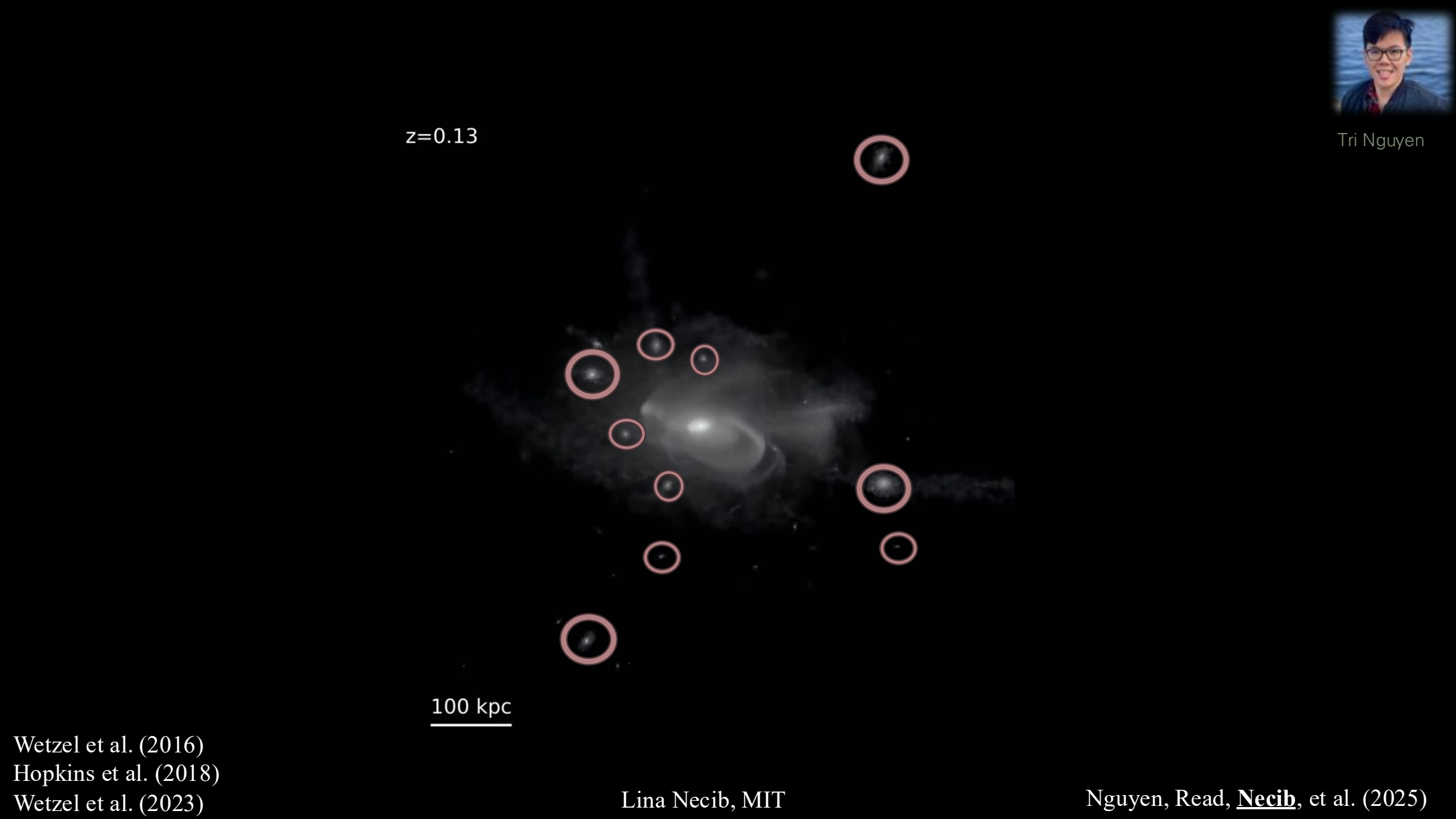


These are idealized simulations though

Cusp: $\gamma = 1$
Core: $\gamma = 0$

Lina Necib, MIT

Nguyen, Mishra-Sharma, Williams, & **Necib** (2022)



Tri Nguyen

Wetzel et al. (2016)
Hopkins et al. (2018)
Wetzel et al. (2023)

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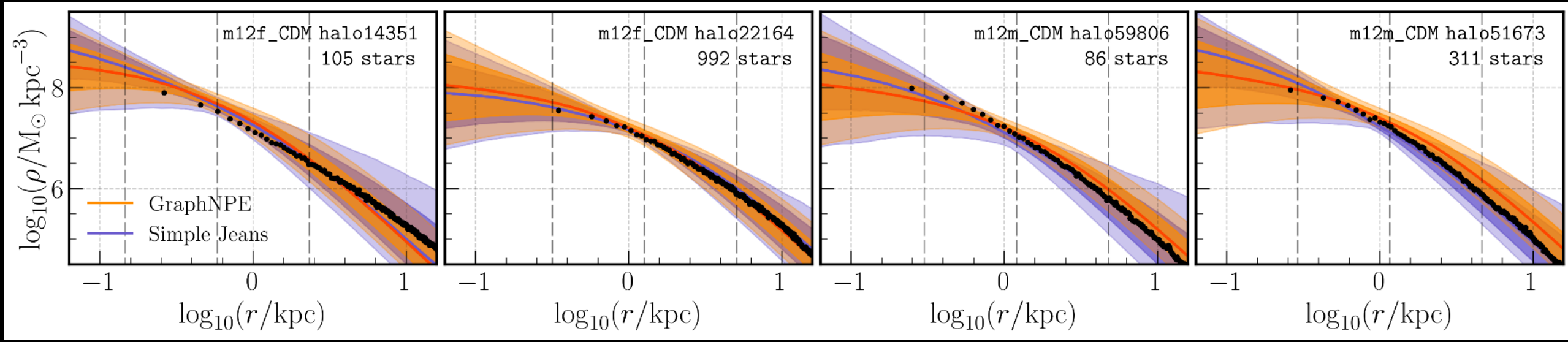
Nguyen, Read, **Necib**, et al. (2025)

APPLYING THIS TO FIRE-2 CDM SATELLITES YIELDS GOOD FITS

Tri Nguyen

GraphNPE

Jeans



Lina Necib, MIT

Strigari et al. (2008)
Nguyen, Read, **Necib**, et al. (2025)

APPLYING THIS TO FIRE-2 SIDM SATELLITES YIELDS GOOD FITS

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GraphNPE

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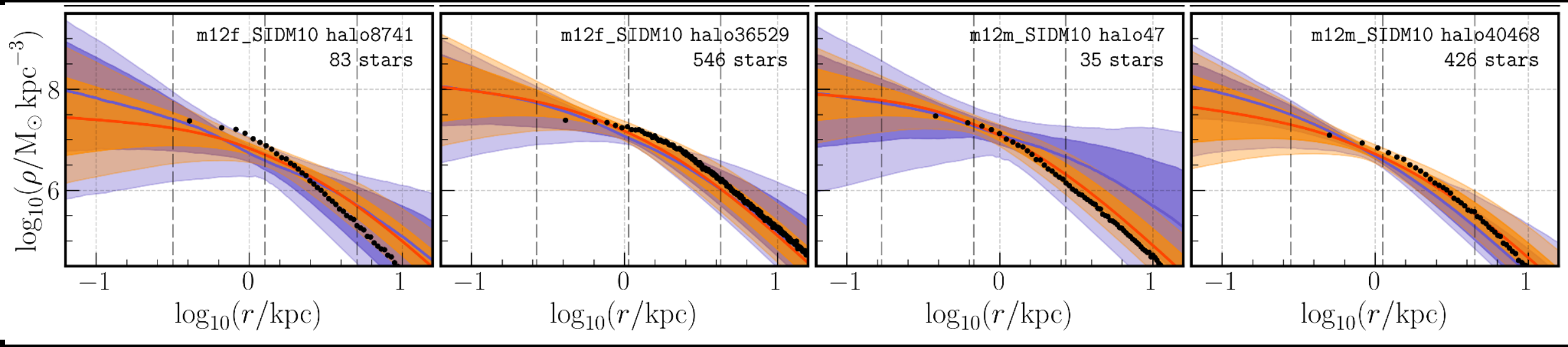
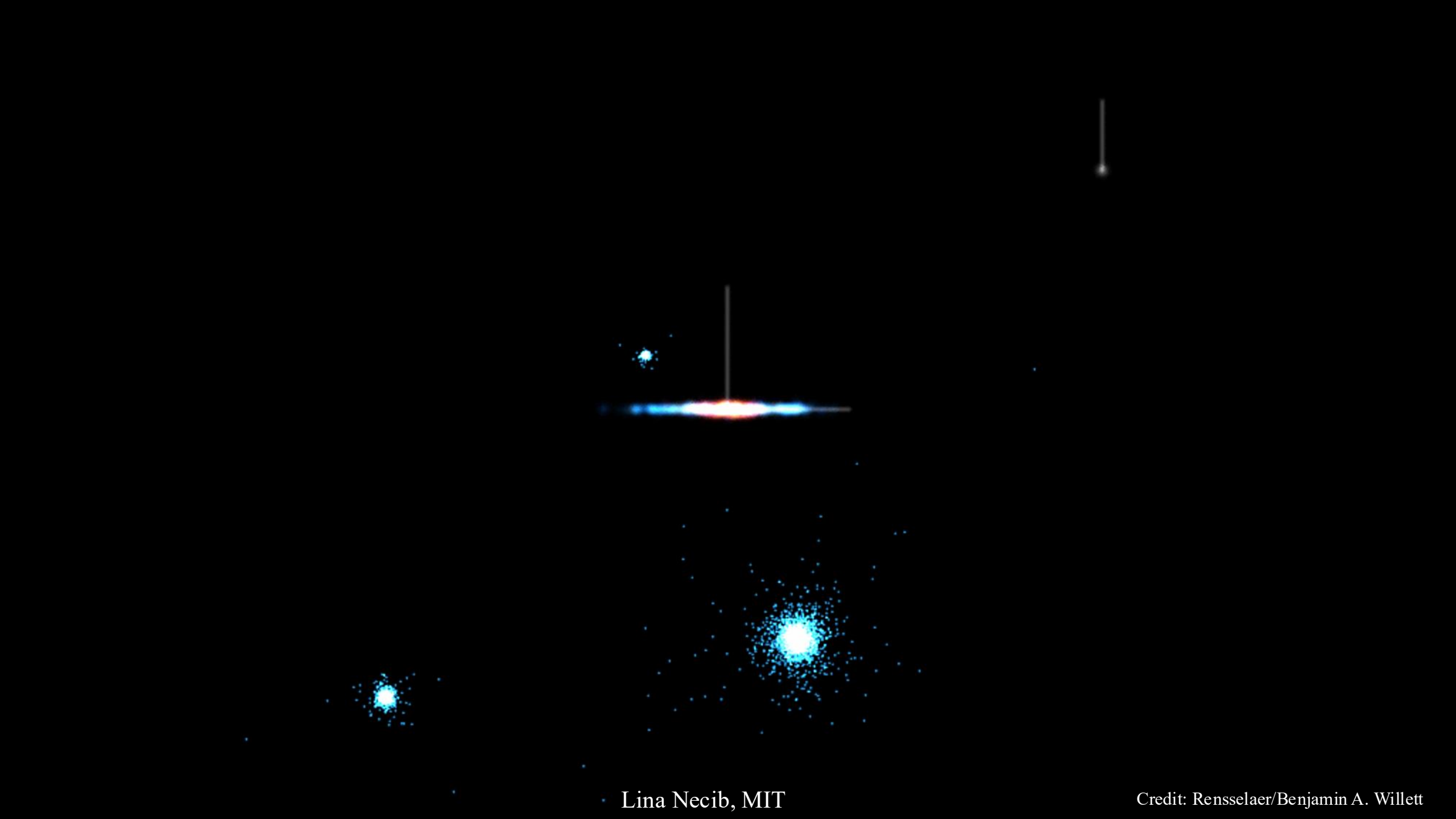


Lina Necib, MIT

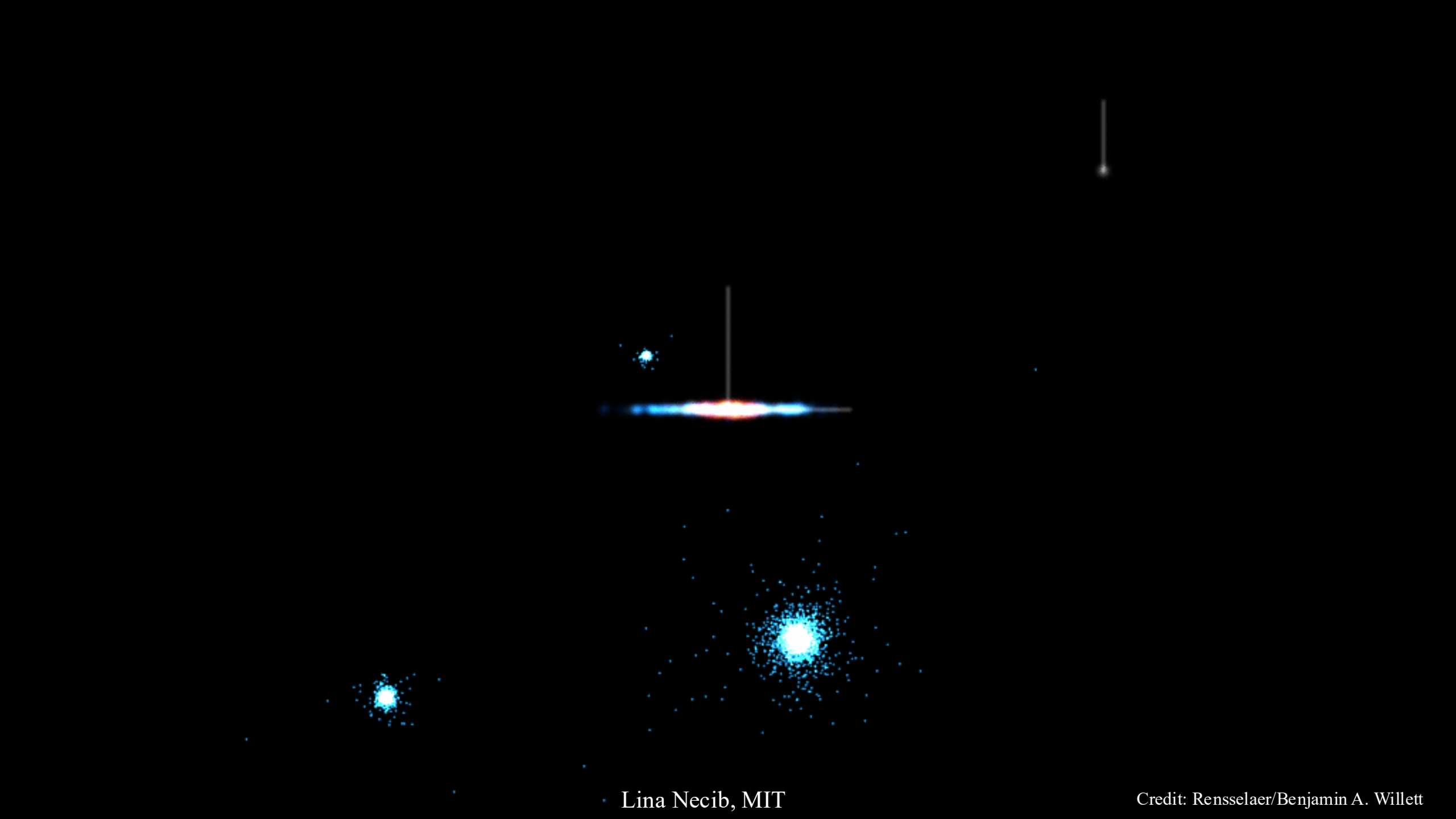
Strigari et al. (2008)

Nguyen, Read, **Necib**, et al. (2025)



Lina Necib, MIT

Credit: Rensselaer/Benjamin A. Willett



Lina Necib, MIT

Credit: Rensselaer/Benjamin A. Willett

z=0.13

100 kpc

Lina Necib, MIT

Shipp, Panithanpaisal, **Necib**, et al. (2023)
Shipp, Riley, …, **Necib**, et al. (2024)

Nora Shipp

z=0.13

100 kpc

Credit: Hernan Stockebrand

Lina Necib, MIT

Shipp, Panithanpaisal, **Necib**, et al. (2023)
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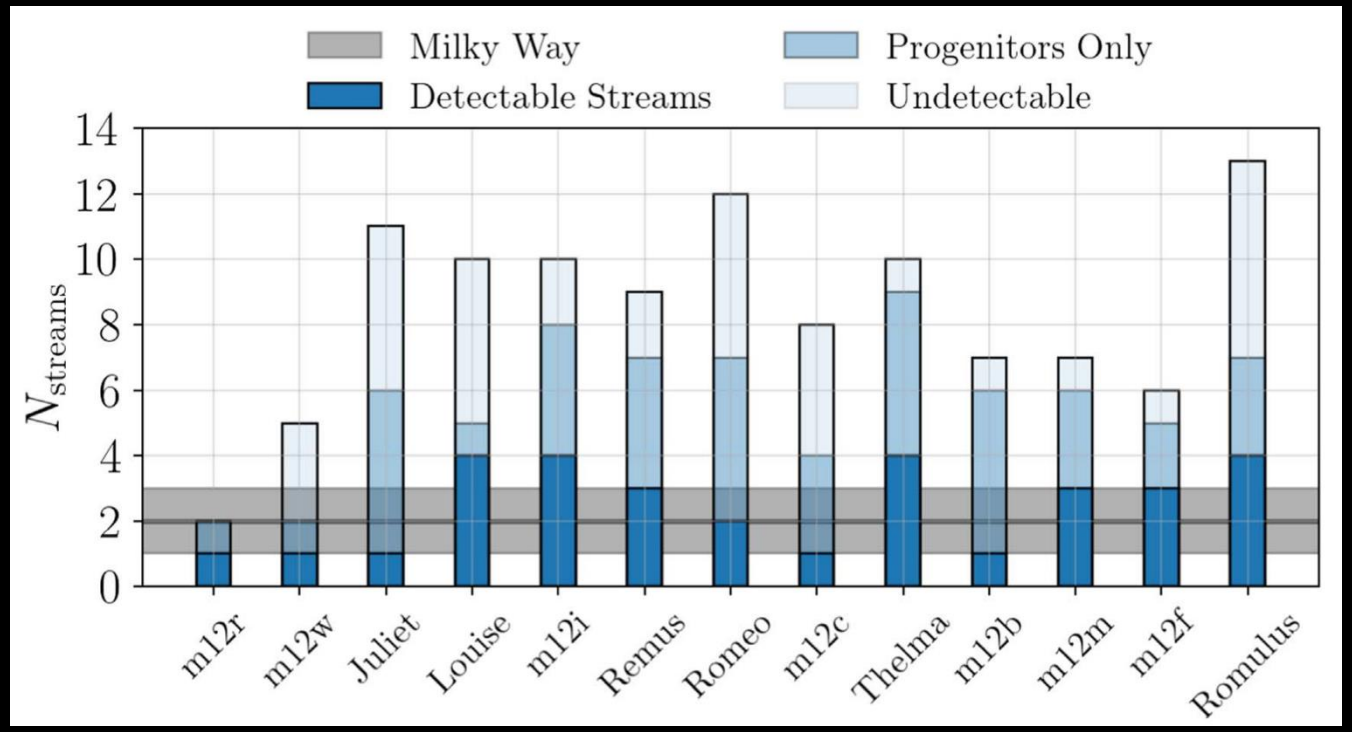
Credit: Hernan Stockebrand

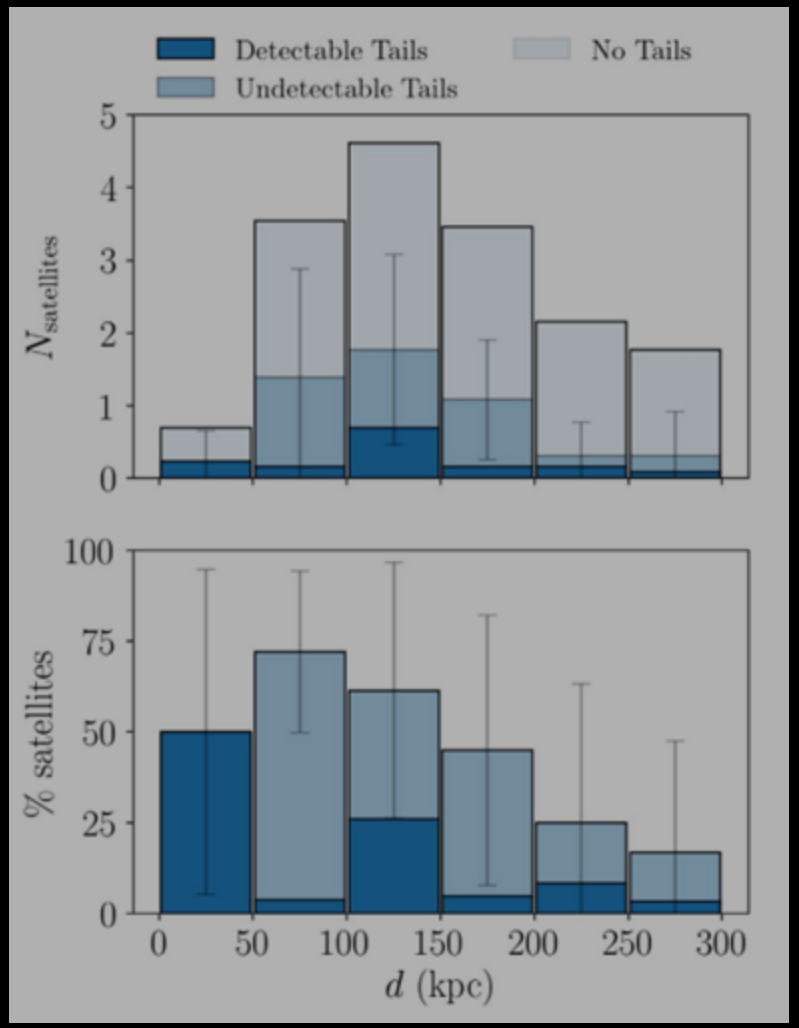
Lina Necib, MIT

Shipp, Panithanpaisal, **Necib**, et al. (2023)
Shipp, Riley, …, **Necib**, et al. (2024)

MOST SIMULATED MILKY WAYS CONTAIN STELLAR STREAMS WHERE WE WOULD ONLY DETECT THE PROGENITOR


Nora Shipp



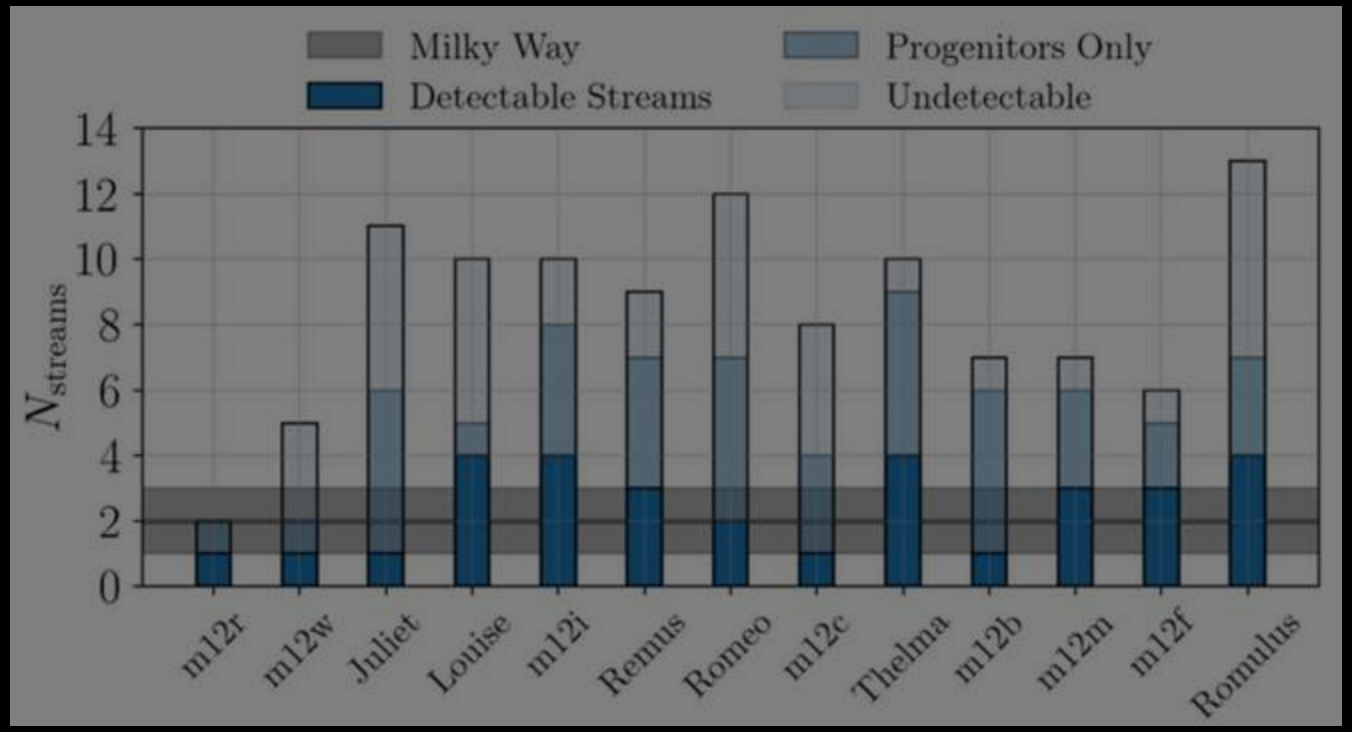


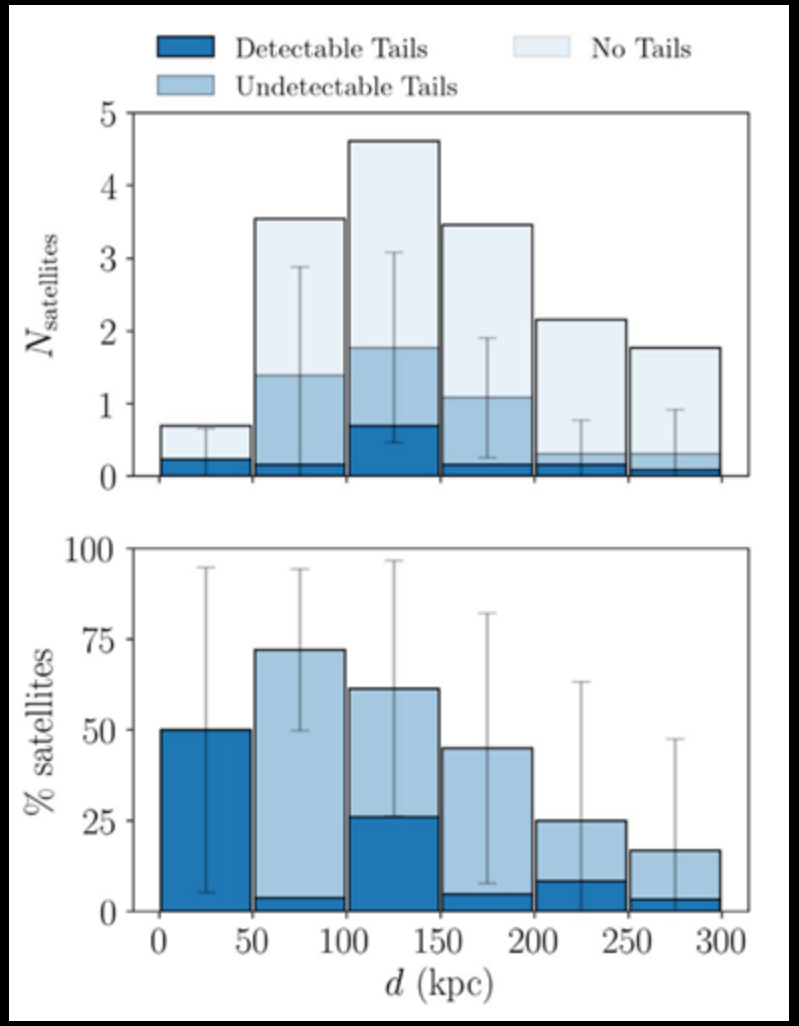
Lina Necib, MIT

Shipp, Panithanpaisal, **Necib**, et al. (2023)
Shipp, Riley, …, **Necib**, et al. (2024)

WE EXPECT THAT MANY OF THE CURRENT SATELLITES OF THE MILKY WAY ARE BEING TIDALLY DISRUPTED.

Nora Shipp





Lina Necib, MIT

Shipp, Panithanpaisal, **Necib**, et al. (2023)
Shipp, Riley, …, **Necib**, et al. (2024)

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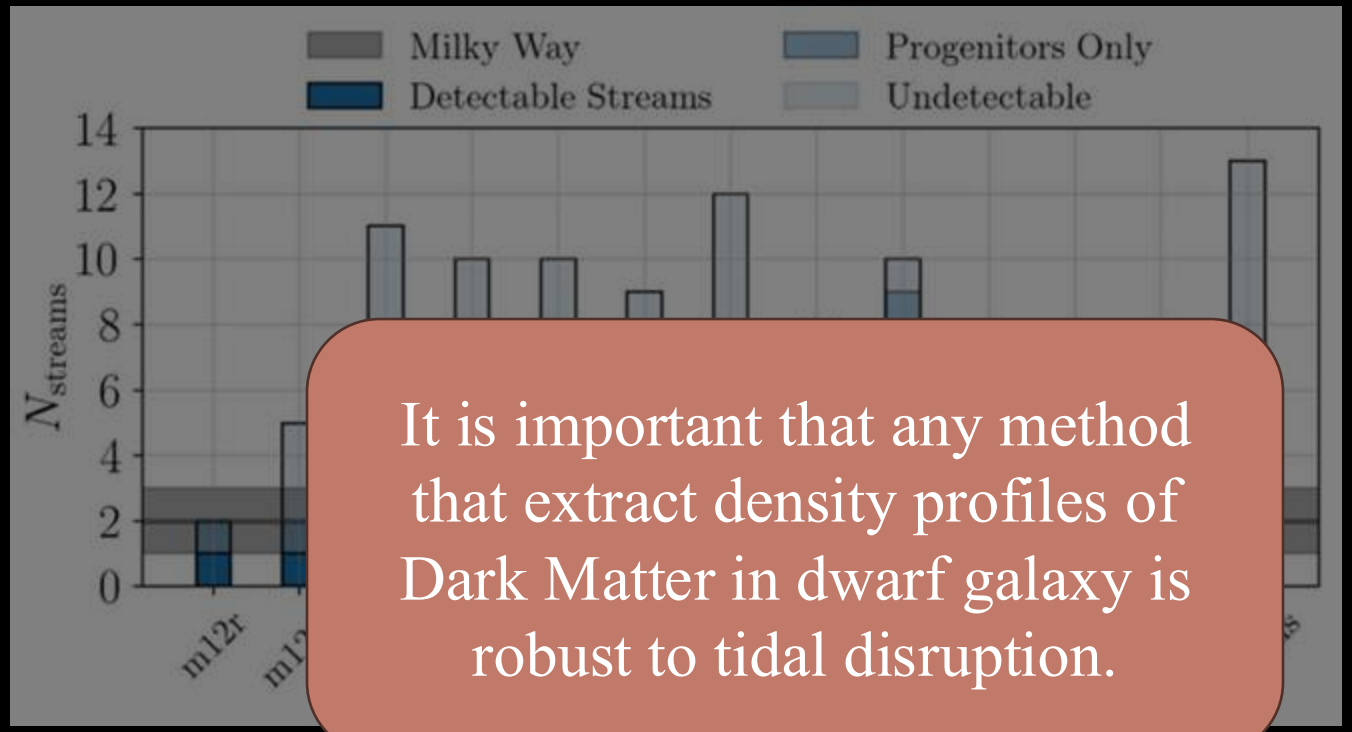
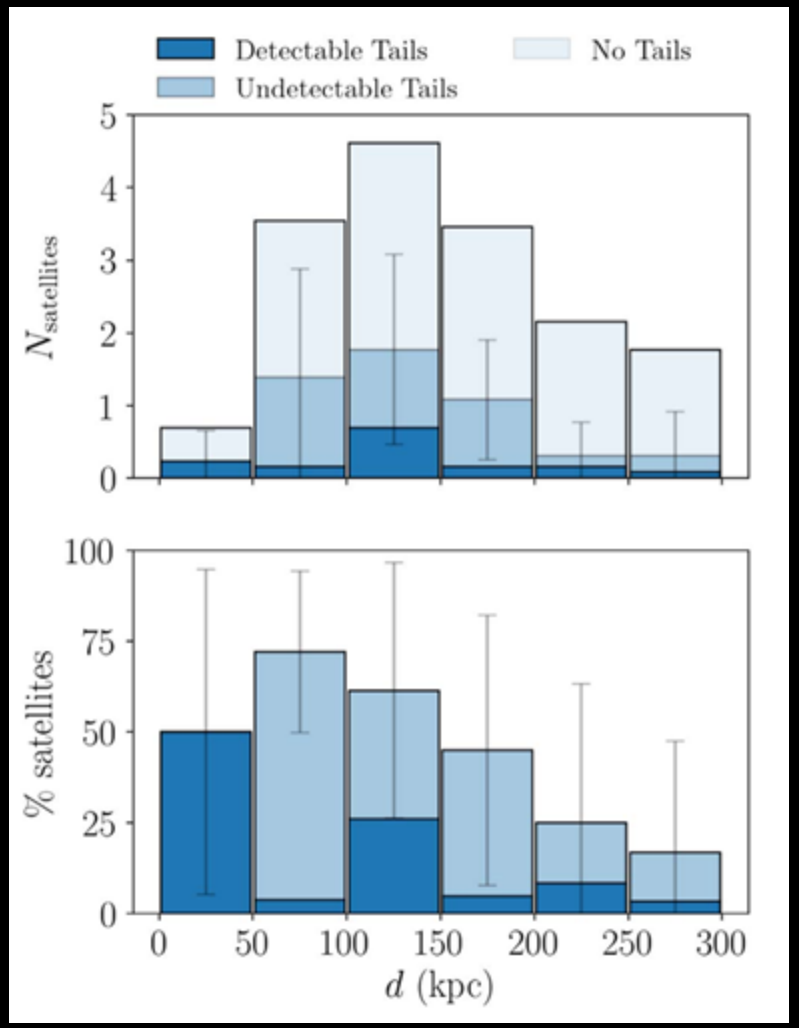


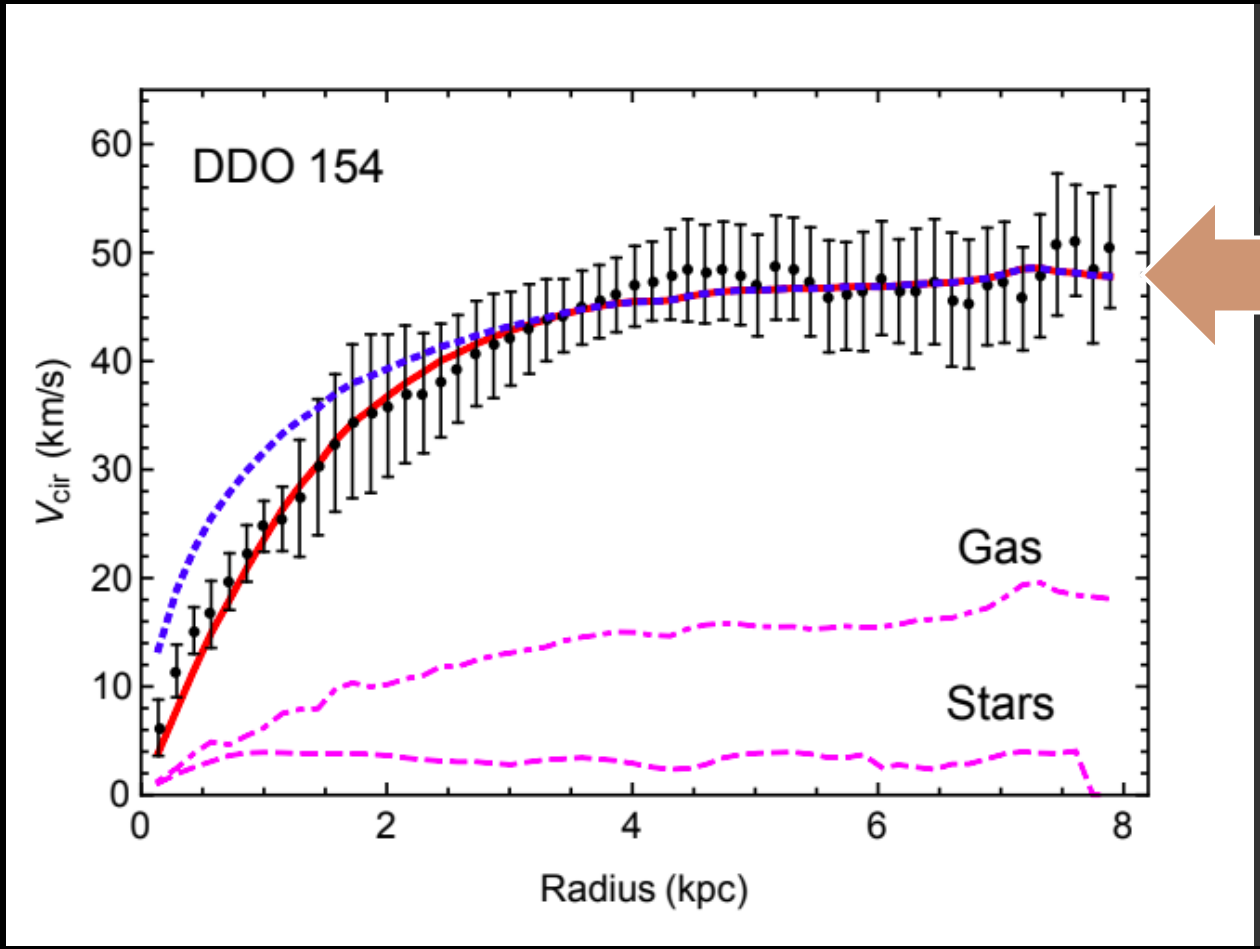
It is important that any method that extract density profiles of Dark Matter in dwarf galaxy is robust to tidal disruption.



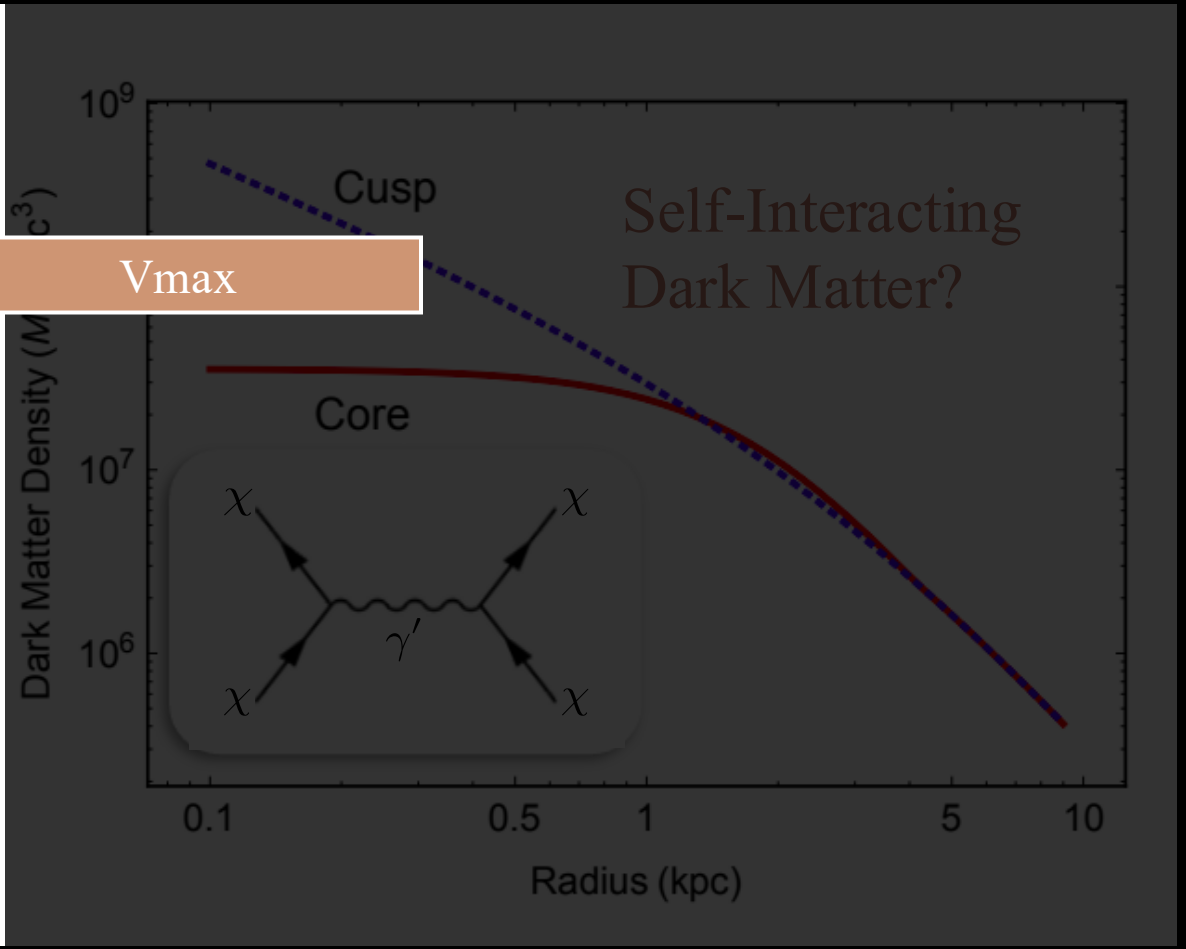
Lina Necib, MIT

Shipp, Panithanpaisal, **Necib**, et al. (2023)
Shipp, Riley, …, **Necib**, et al. (2024)

EARLY WORKS HAVE SHOWN THAT DWARF GALAXIES MIGHT HAVE CORES



Vmax



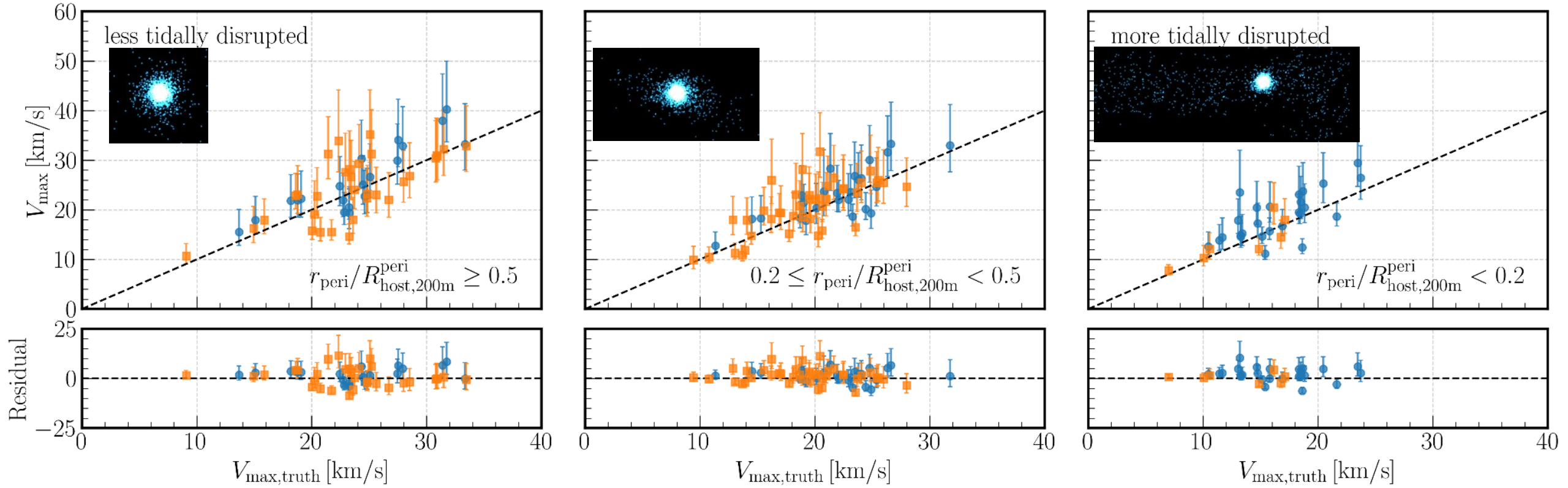
Spergel & Steinhardt (2000)

Oh et al. (2015)
Tulin & Yu (2017)

WE OBTAIN A GOOD FIT FOR PROPERTIES OF GALAXIES EVEN IN THE CASES OF TIDAL DISRUPTION: VMAX

Tri Nguyen

SIDM CDM



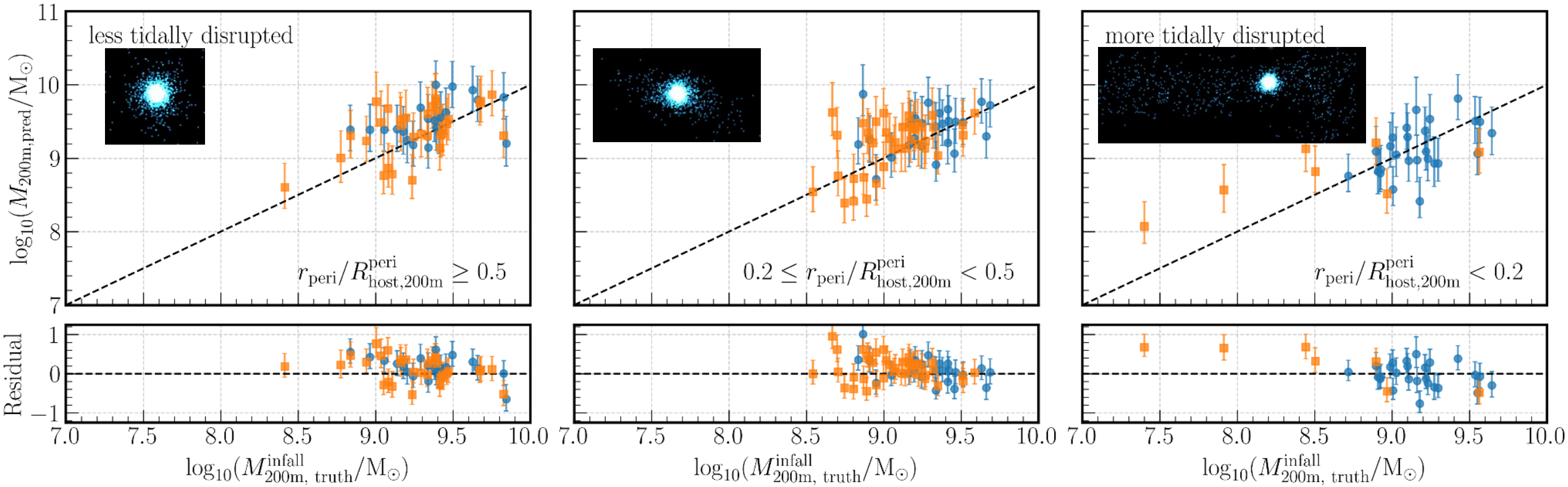
Lina Necib, MIT

Nguyen, Read, **Necib**, et al. (2025)

WE OBTAIN A GOOD FIT FOR PROPERTIES OF GALAXIES EVEN IN THE CASES OF TIDAL DISRUPTION: M200

Tri Nguyen

SIDM CDM



Lina Necib, MIT

Nguyen, Read, **Necib**, et al. (2025)

WE OBTAIN A GOOD FIT FOR PROPERTIES OF GALAXIES EVEN IN THE CASES OF TIDAL DISRUPTION: M200

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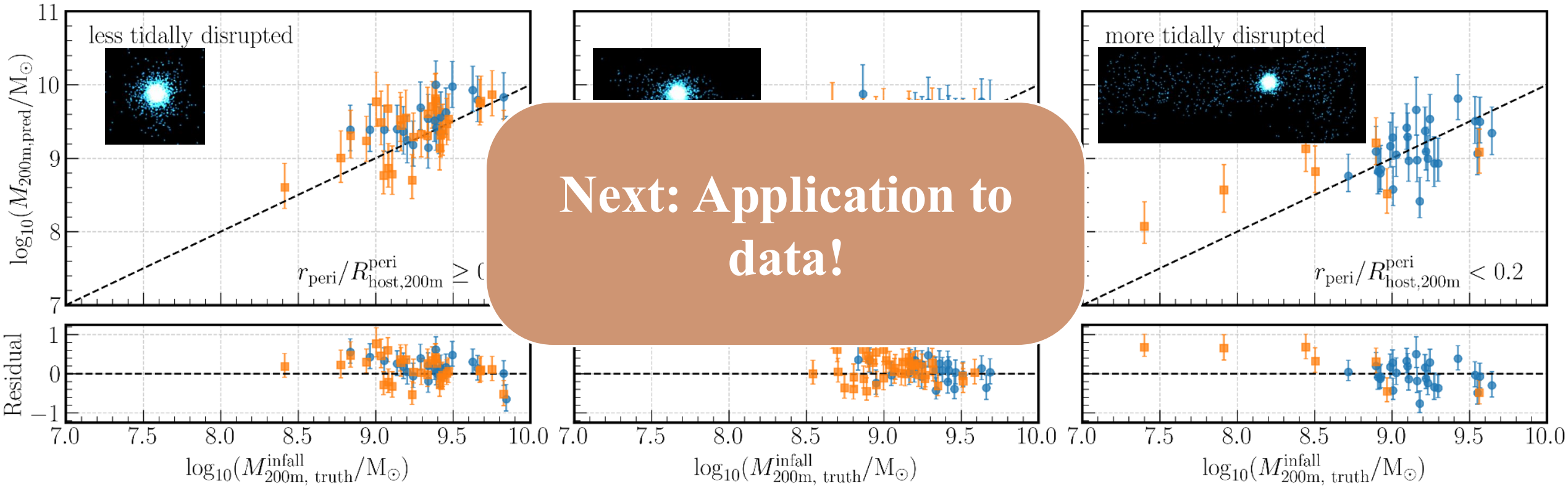
SIDM CDM



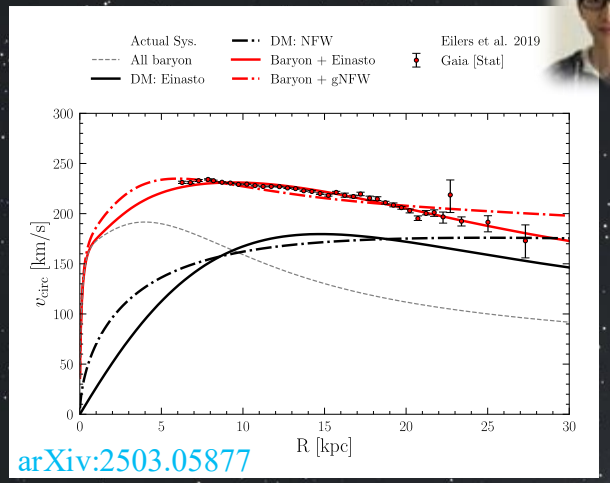
Next: Application to data!

Lina Necib, MIT

Nguyen, Read, **Necib**, et al. (2025)

TODAY'S TALK: MAPPING THE DARK MATTER DENSITY PROFILE IN THE MILKY WAY AND DWARF GALAXIES

Xiaowei Ou

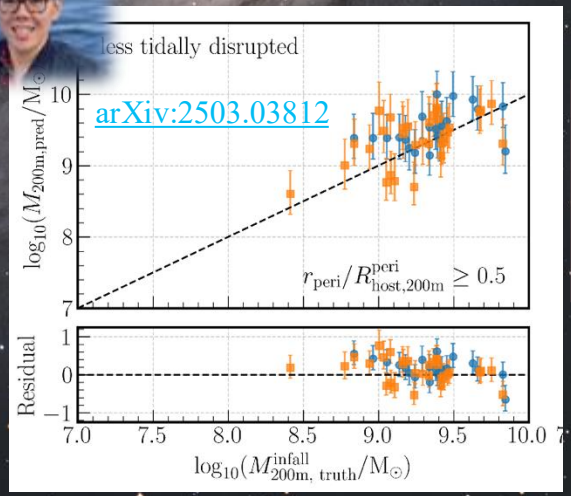


arXiv:2503.05877

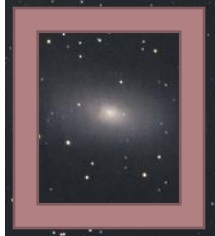
Used synthetic surveys to build the true systematic uncertainties of the circular velocity measurements of the Milky Way, and obtained a new Dark Matter density profile.

Galactic Center/Halo

Tri Nguyen

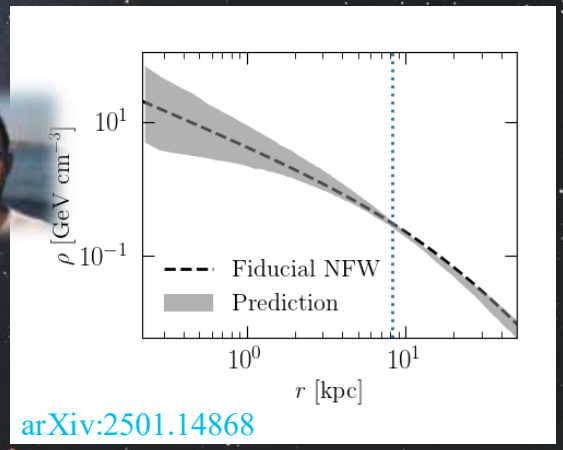


Dwarf Galaxies



Used Graph Neural Networks to construct the density profiles of Dark Matter in dwarf galaxies, based on line-of-sight velocities.
Tested this new method on FIRE simulations to show robustness. Next we will apply this to data.

Abdelaziz Hussein



arXiv:2501.14868

Built a theoretical range for the density profile of Dark Matter at the center of the Milky Way, based on adiabatic contraction and baryonic feedback.

GOAL:
MAPPING FROM THE STARS TO THE DARK MATTER

Thank you!