GUT-less SUSY Phenomenology

Pearl Sandick

University of Minnesota

Ellis, Olive & PS, Phys. Lett. B **642** (2006) 389
Ellis, Olive & PS, JHEP 06 (2007) 079

Why we like SUSY

- Solves the Naturalness Problem
- Gauge coupling unification (GUTs)
- Predicts a light Higgs boson

R-Parity conservation →

What We Do

- SUSY must be broken, so introduce soft SUSY-breaking parameters and assume high (GUT) scale values for them
- Evolve parameters down to weak scale using RGEs of low energy effective theory (MSSM)
- CMSSM: GUT-scale universality of soft breaking parameters
  - 5 inputs: m_0, $m_{1/2}$, A_0, $\tan(\beta)$, $\mathrm{sign}(\mu)$

Some other scale?

GUT-less CMSSM

- Assume unification of soft SUSY-breaking parameters at some $M_{in} < M_{GUT}$
  - Constraints from colliders and cosmology:

$m_h > 114$ GeV	LEP
$m_{\chi^\pm} > 104$ GeV	LEP
BR($b \to s\gamma$)	HFAG
BR($B_s \to \mu^+\mu^-$)	CDF
$(g_\mu - 2)/2$	g-2 collab.

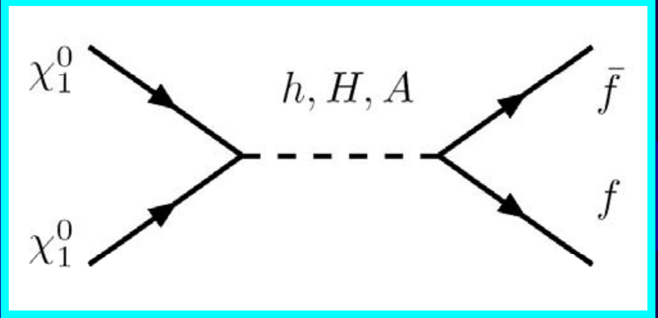
$0.09 \leq \Omega_\chi h^2 \leq 0.12$

SUSY Dark Matter

- Solve Boltzmann rate equation:

$$\frac{dn_\chi}{dt} = -3Hn_\chi - \langle \sigma v_{rel} \rangle \left[n_\chi^2 - (n_\chi^{eq})^2 \right]$$

- Special Situations:
  - s - channel poles
    - 2 $m_\chi \approx m_A$
  - thresholds
    - 2 $m_\chi \approx$ final state mass
  - Coannihilations
    - $m_\chi \approx m_{\text{other sparticle}}$



Evolution of the Soft Mass Parameters

- First look at gaugino and scalar mass evolution.

Gauginos (1-Loop):

$$M_a(Q) = \frac{\alpha_a(Q)}{\alpha_a(M_{GUT})} M_a(M_{GUT}) \Rightarrow M_a(Q) = \frac{\alpha_a(Q)}{\alpha_a(M_{in})} m_{1/2}$$

Running of gauge couplings identical to CMSSM case, so low scale gaugino masses are all closer to $m_{1/2}$ as M_{in} is lowered.

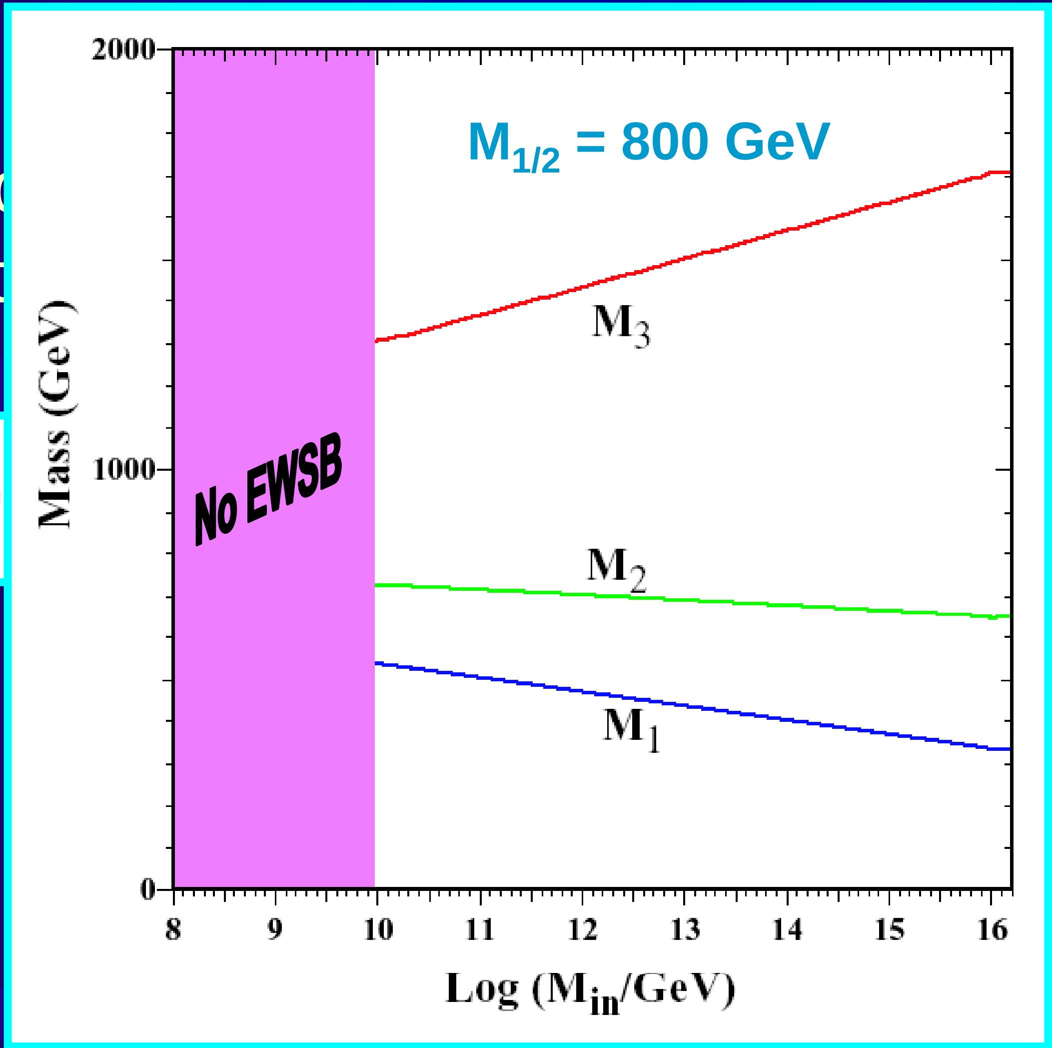
Evolution of the Soft Mass Parameters

- First lo
  evolut

  Gaugi

$M_a(Q)$ ... $\frac{(Q)}{M_{in})} m_{1/2}$

al to
no
is



Evolution of the Soft Mass Parameters

- First look at gaugino and scalar mass evolution.

Scalars (1-Loop):

$$m_{0_i}^2(Q) = m_0^2 + C_i(Q, M_{GUT})\, m_{1/2}^2$$

$$\Rightarrow \quad m_{0_i}^2(Q) = m_0^2(M_{in}) + C_i(Q, M_{in}) m_{1/2}^2$$

$$C_i \to 0 \text{ as } M_{in} \to Q$$

As $M_{in} \to$ low scale Q, expect low scale scalar masses to be closer to m_0.

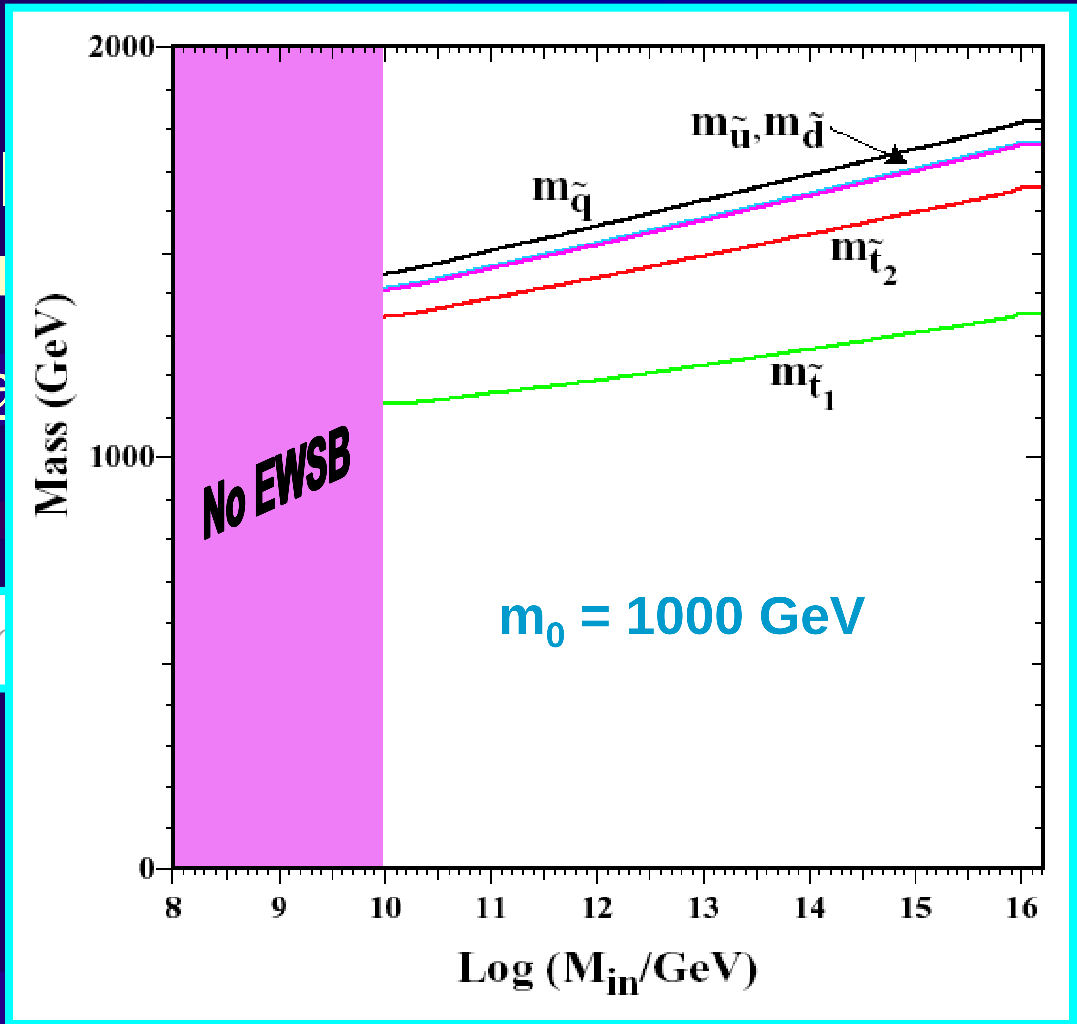
Evolution of the Soft Mass Parameters

- First l
  evolu

Scala

m … $M_{in} \to Q$

…cale



Evolution of the Soft Mass Parameters

- Higgs mass parameter, μ (tree level):

$$\mu^2 = \frac{m_1^2 - m_2^2 \tan^2\beta}{\tan^2\beta - 1} - \frac{M_Z^2}{2}$$

As $M_{in} \rightarrow$ low scale Q, expect low scale scalar masses to be closer to $m_{0.}$

μ^2 becomes generically smaller as M_{in} is lowered.

Mass Evolution with M_{in}

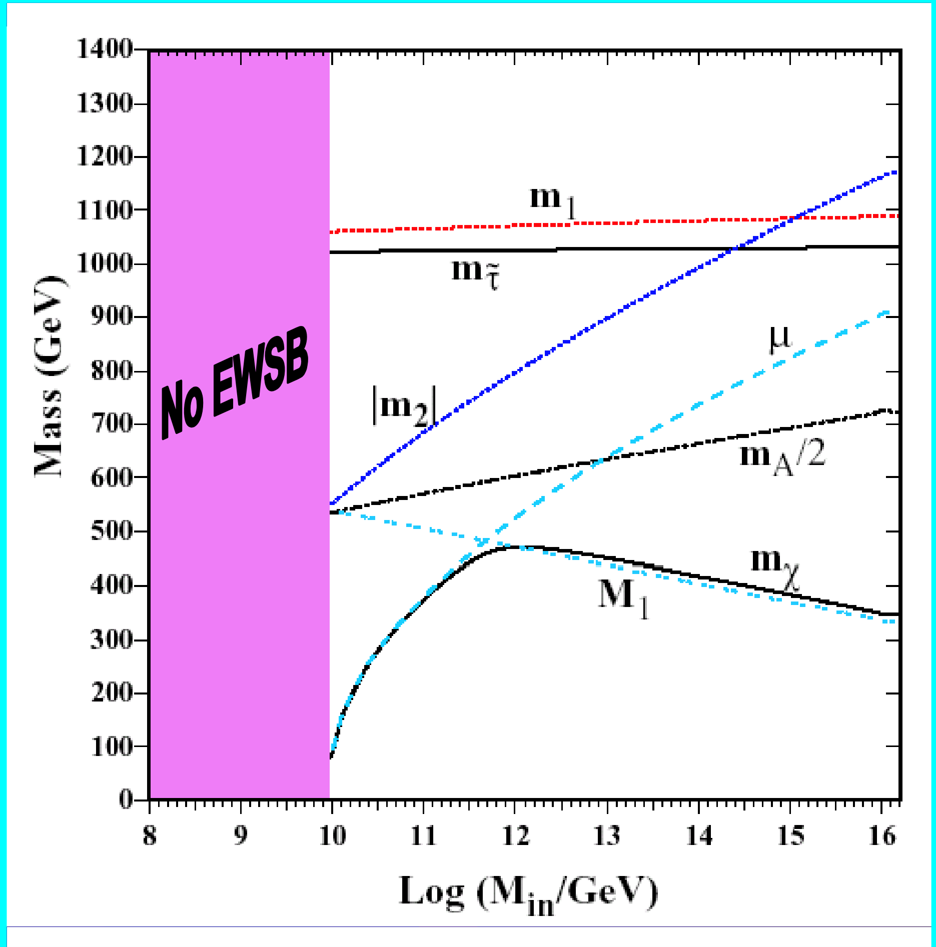
$m_{1/2}$ = 800 GeV

m_0= 1000 GeV

A_0 = 0

tan(β) = 10

μ > 0



How do we expect the constraints to evolve?

- m_A decreases logarithmically with M_{in}
  - BR($b \to s\,\gamma$) and BR($B_s \to \mu^+\mu^-$) at large tan(β) have important contributions from heavy Higgs exchange. These constraints will become more important as M_{in} is lowered.
- μ decreases as M_{in} is lowered.
  - Expect that the unphysical region where $\mu^2 < 0$ encroaches farther into the plane.
  - When the LSP is bino-like, its mass *increases* as M_{in} is lowered, so the forbidden stau LSP region encroaches into the plane. When the LSP becomes Higgsino-like, it's mass *decreases* as M_{in} is lowered, so the stau LSP boundary falls back down.

Neutralinos and Charginos

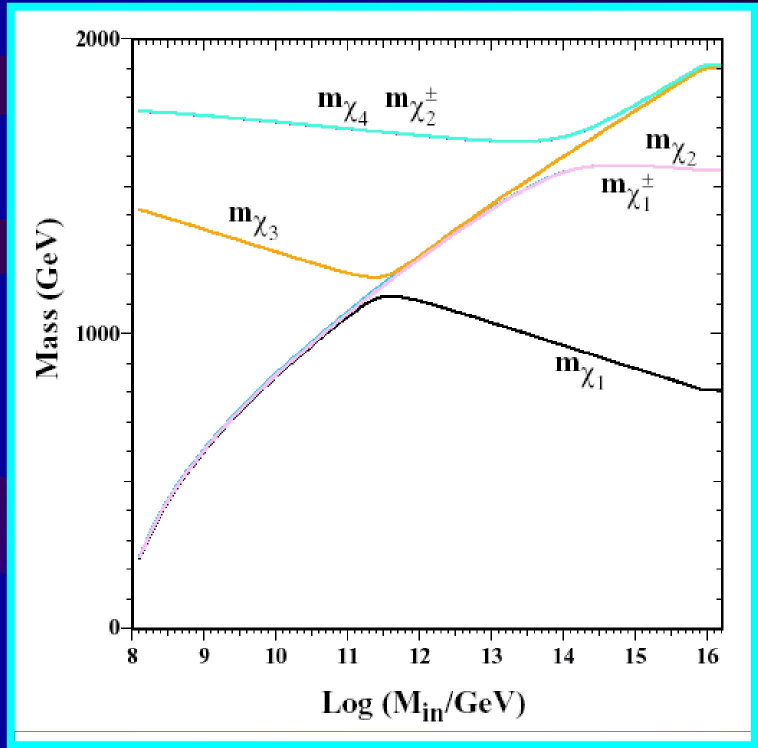
$m_{1/2}$ = 1800 GeV

m_0= 1000 GeV

A_0 = 0

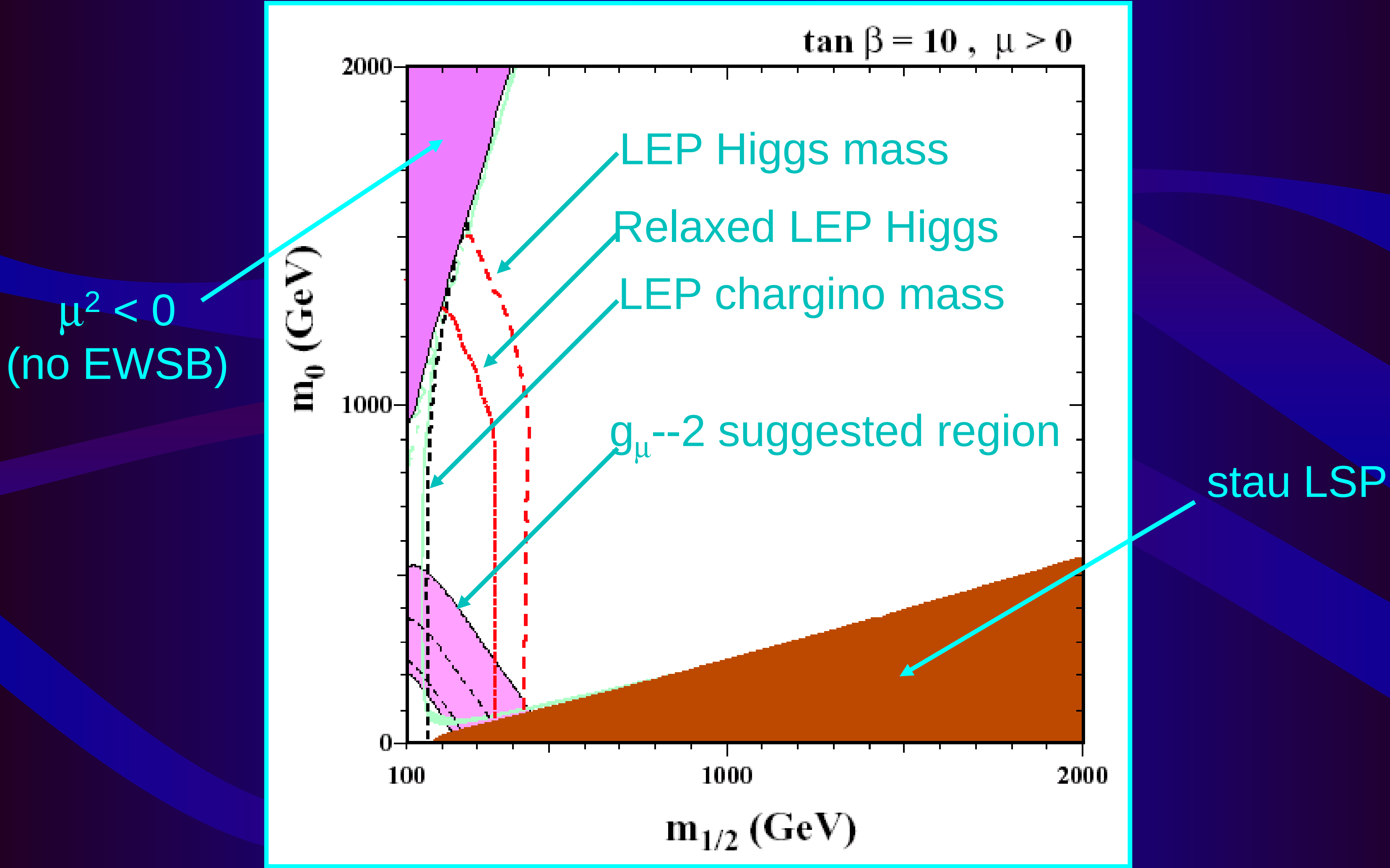
tan(β) = 10

μ > 0

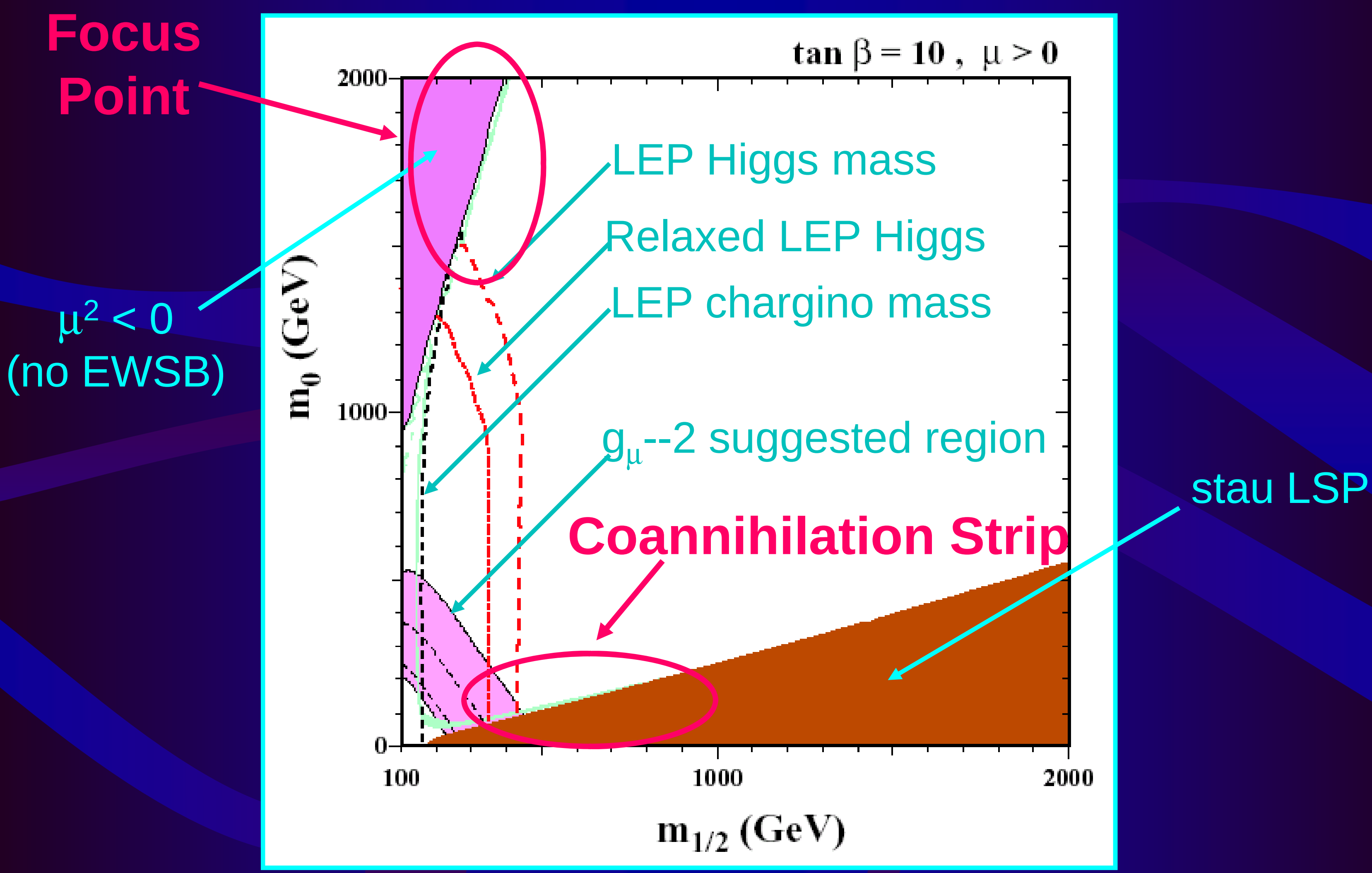


Must properly include coannihilations involving all three lightest neutralinos!

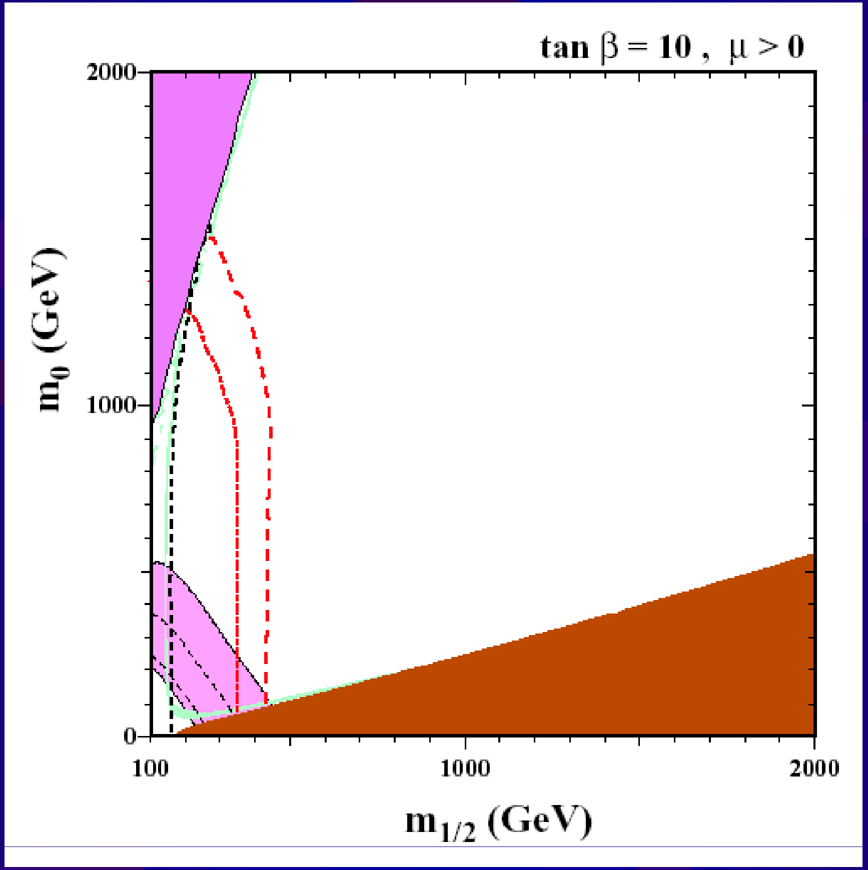
Standard CMSSM



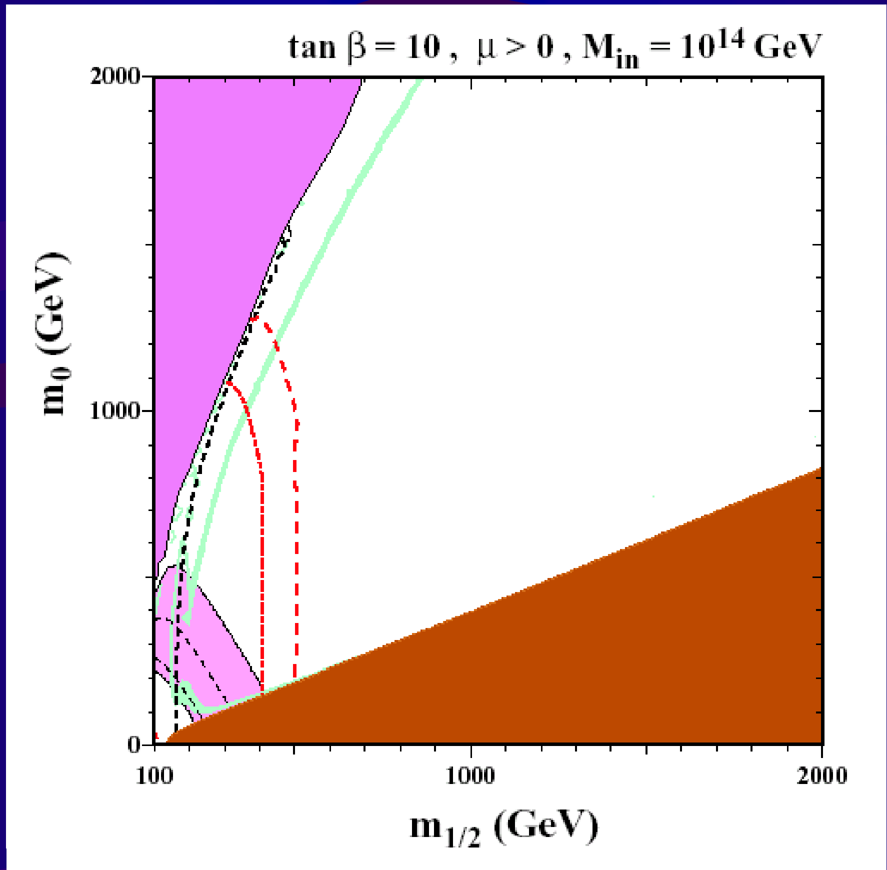
Standard CMSSM



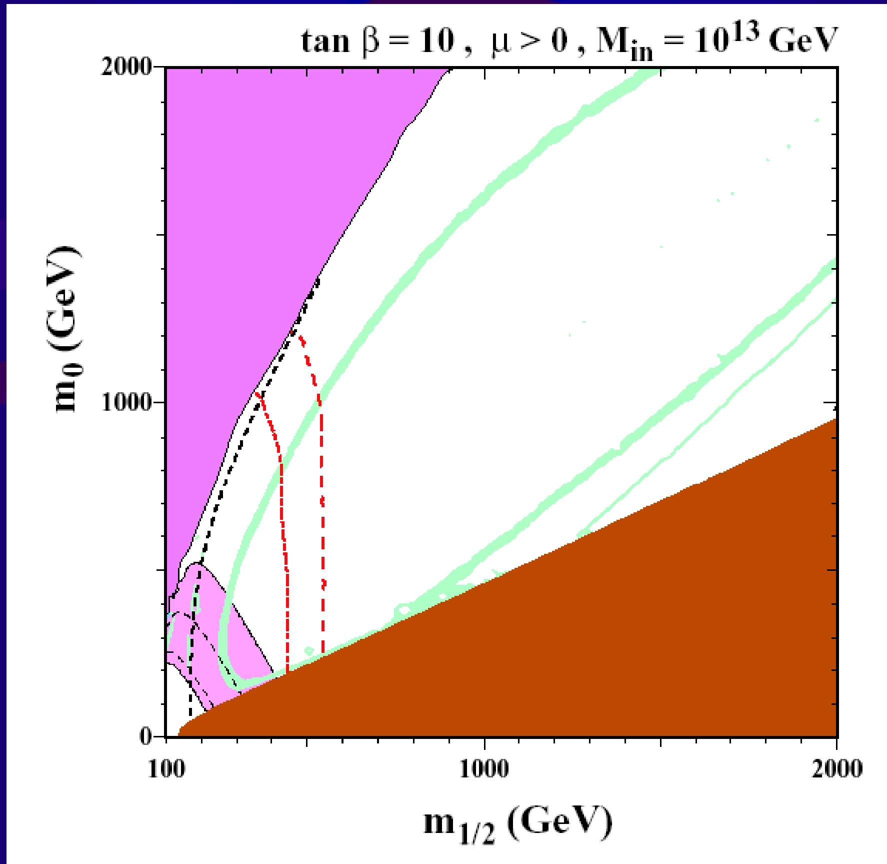
Lowering M_{in} - tan(β) = 10



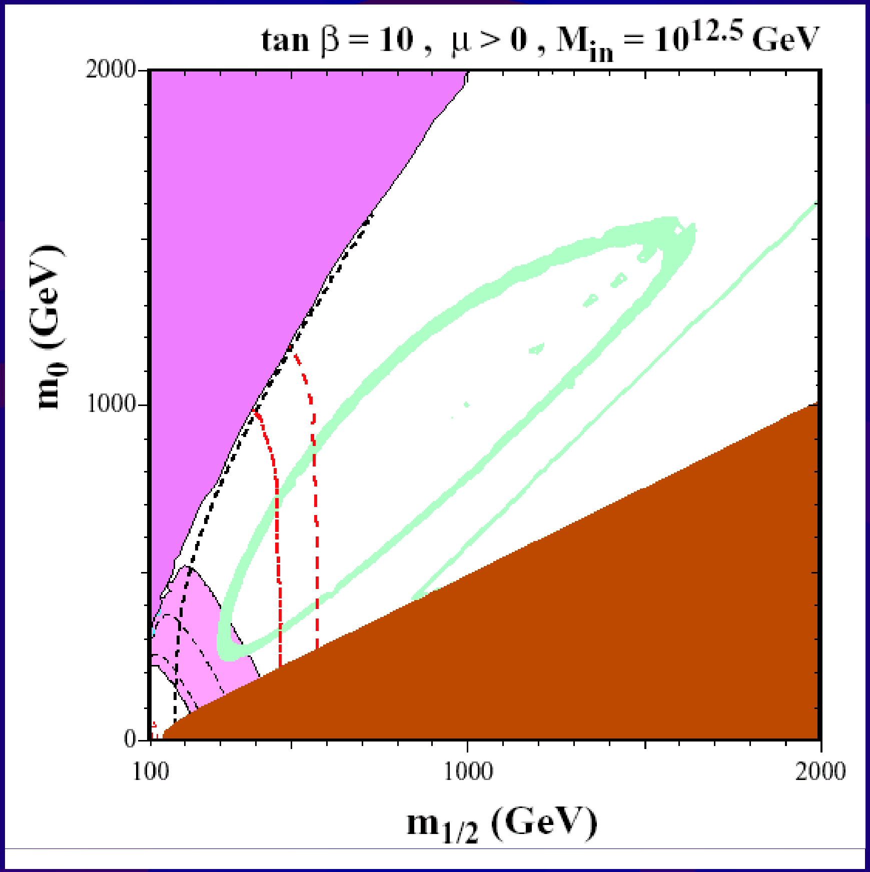
Lowering M_{in} - tan(β) = 10



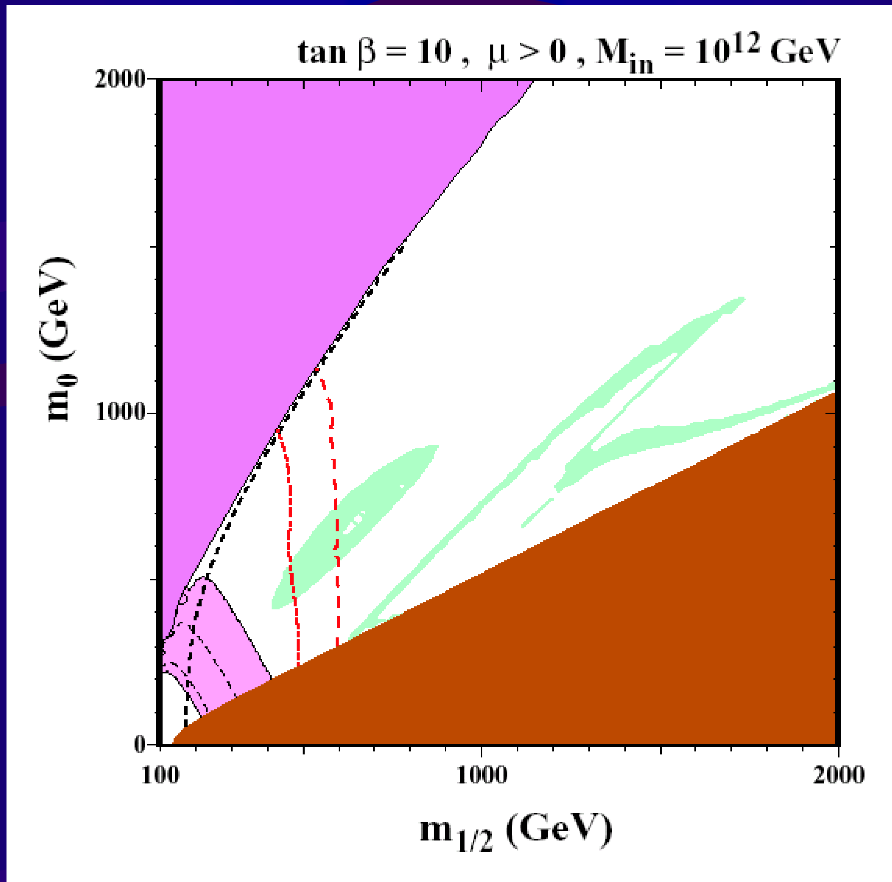
Lowering M_{in} - tan(β) = 10



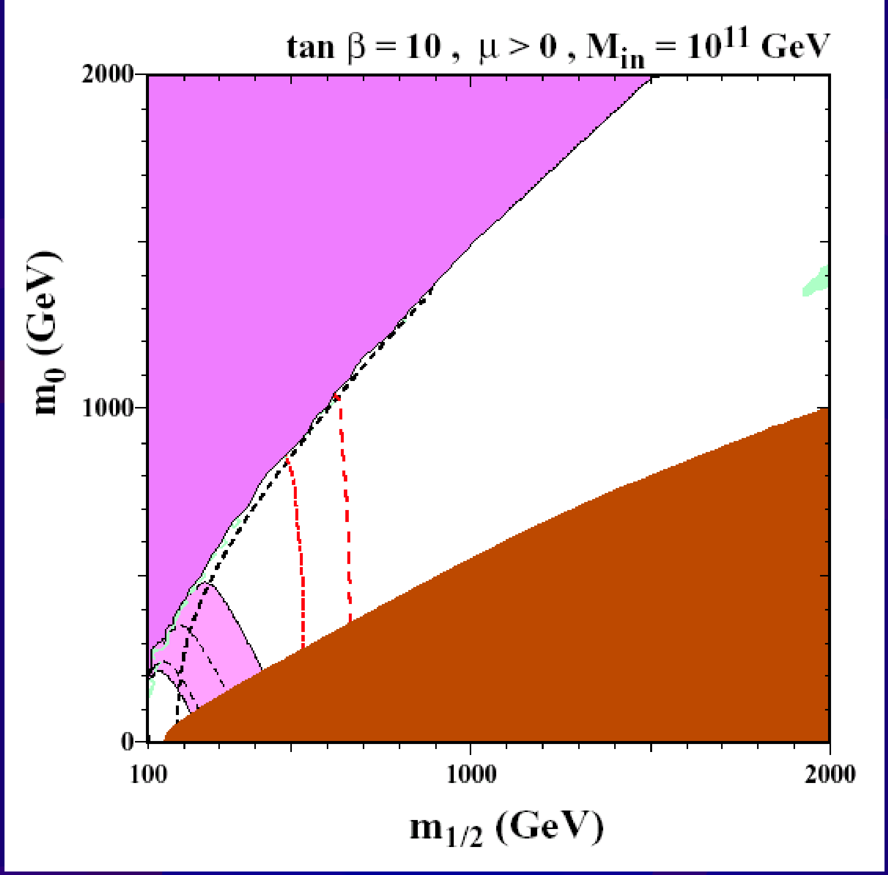
Lowering M_{in} - tan(β) = 10



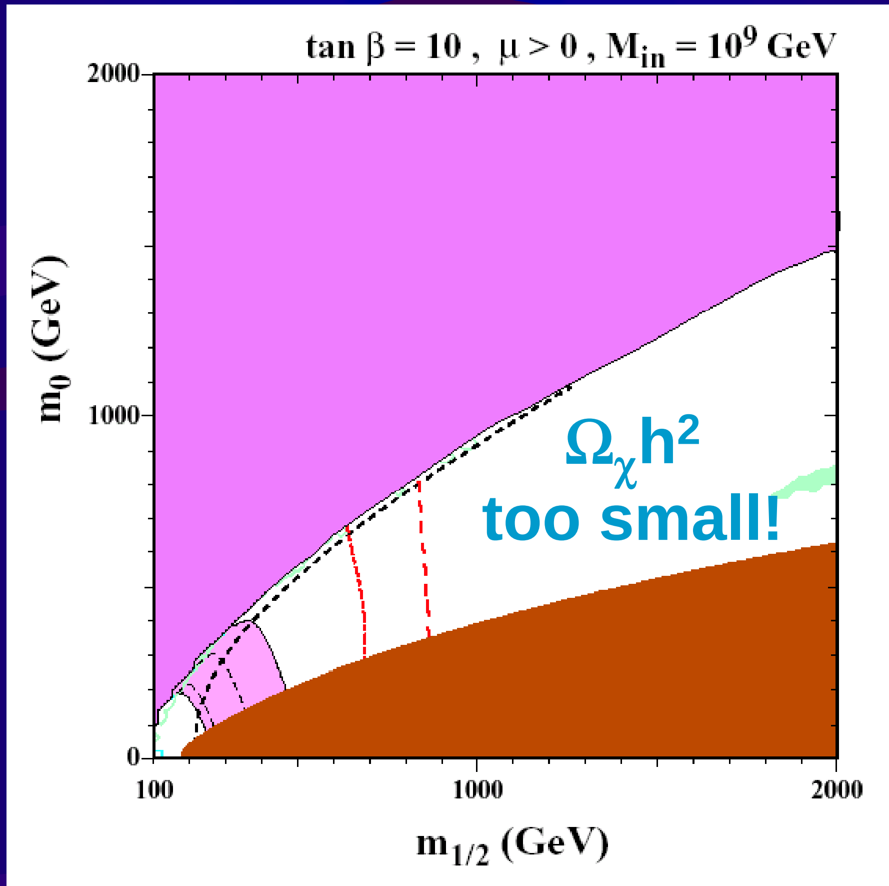
Lowering M_{in} - tan(β) = 10



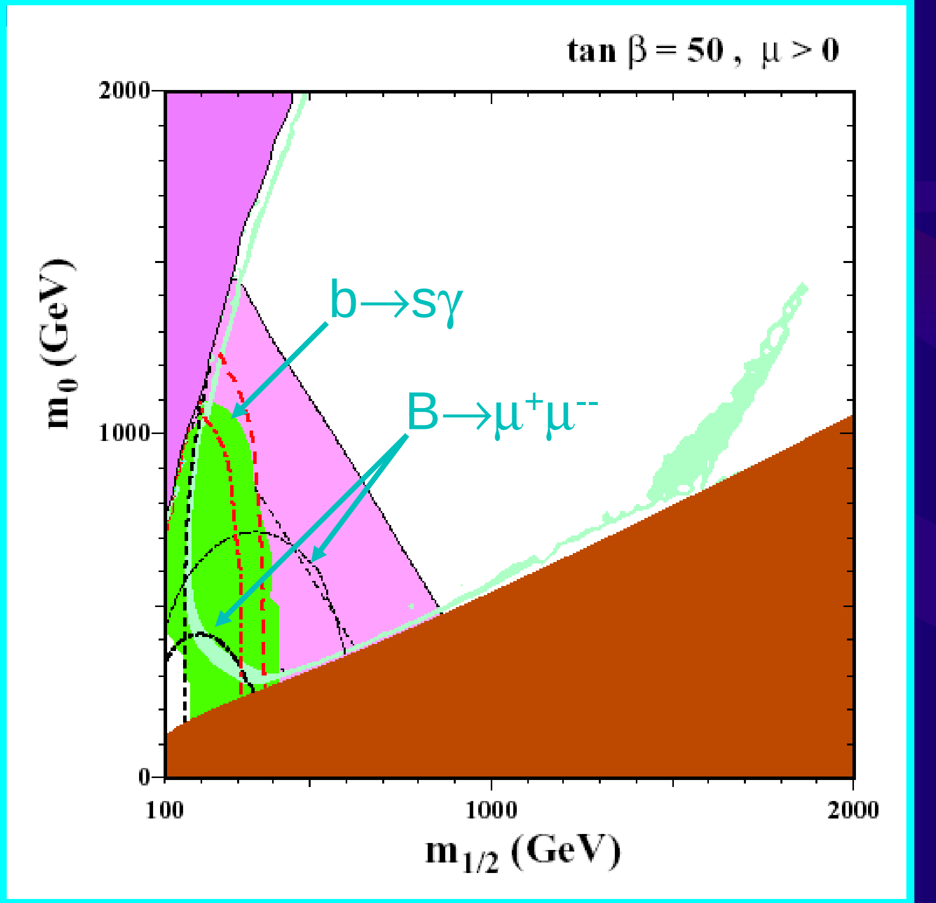
Lowering M_{in} - tan(β) = 10



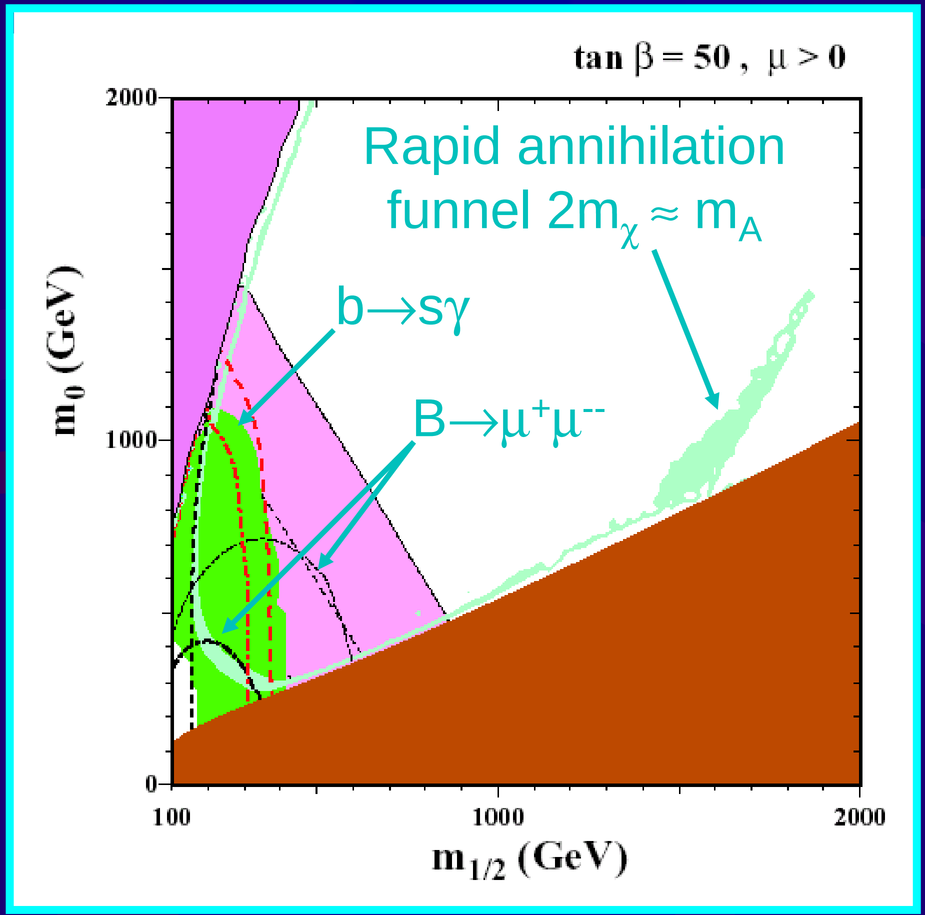
Lowering M_{in} - $\tan(\beta) = 10$



Large tan(β)

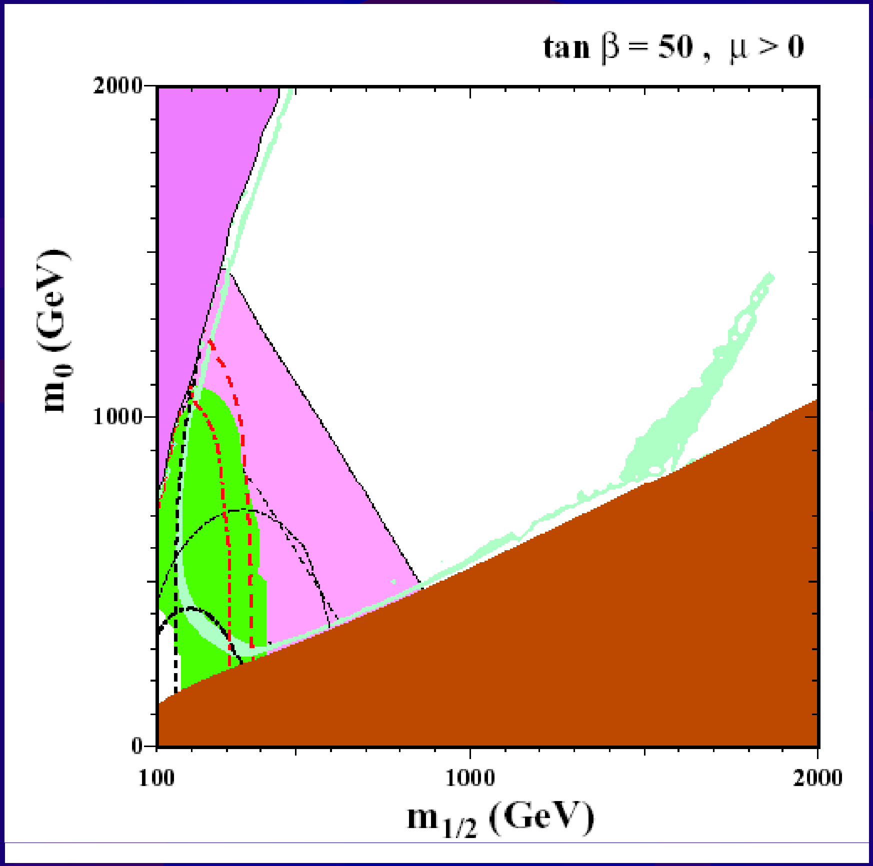


Large tan(β)

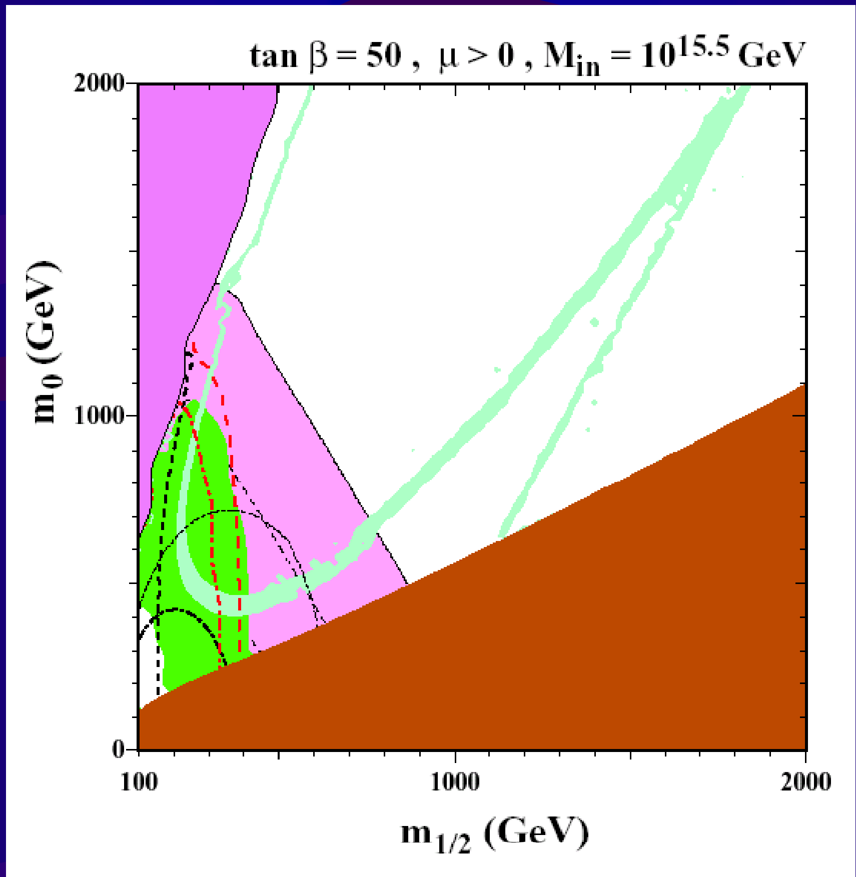


tan β = 50 , μ > 0
Rapid annihilation
funnel 2m$_\chi$ ≈ m$_A$
b→sγ
B→μ$^+$μ^{--}
2000
1000
0
100
1000
2000
m$_0$ (GeV)
m$_{1/2}$ (GeV)

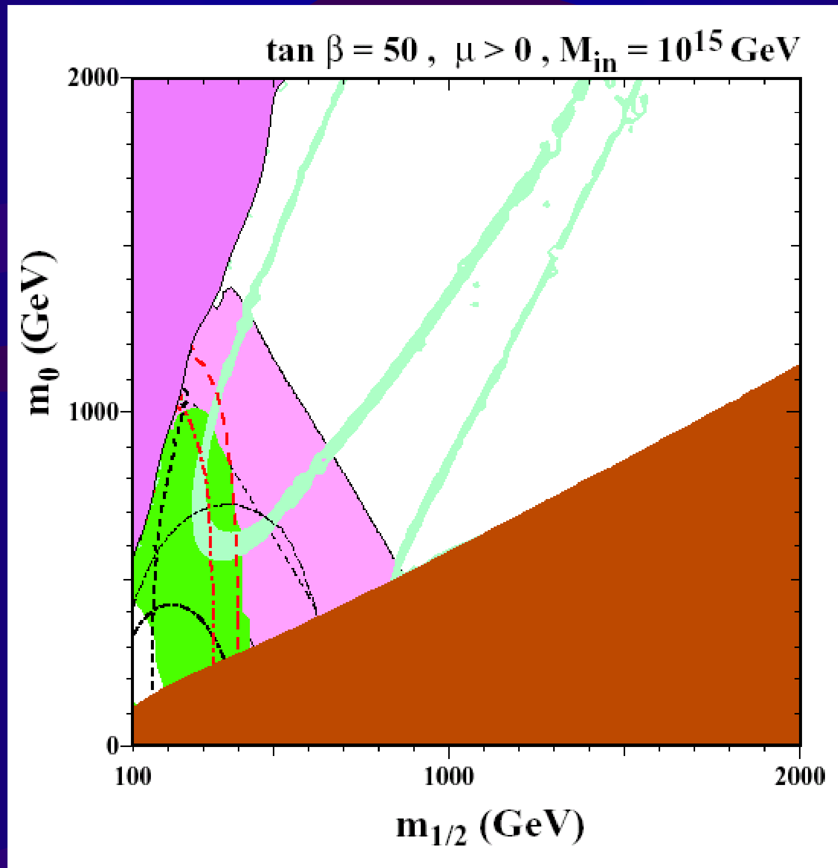
Lowering M_{in} - tan(β) = 50



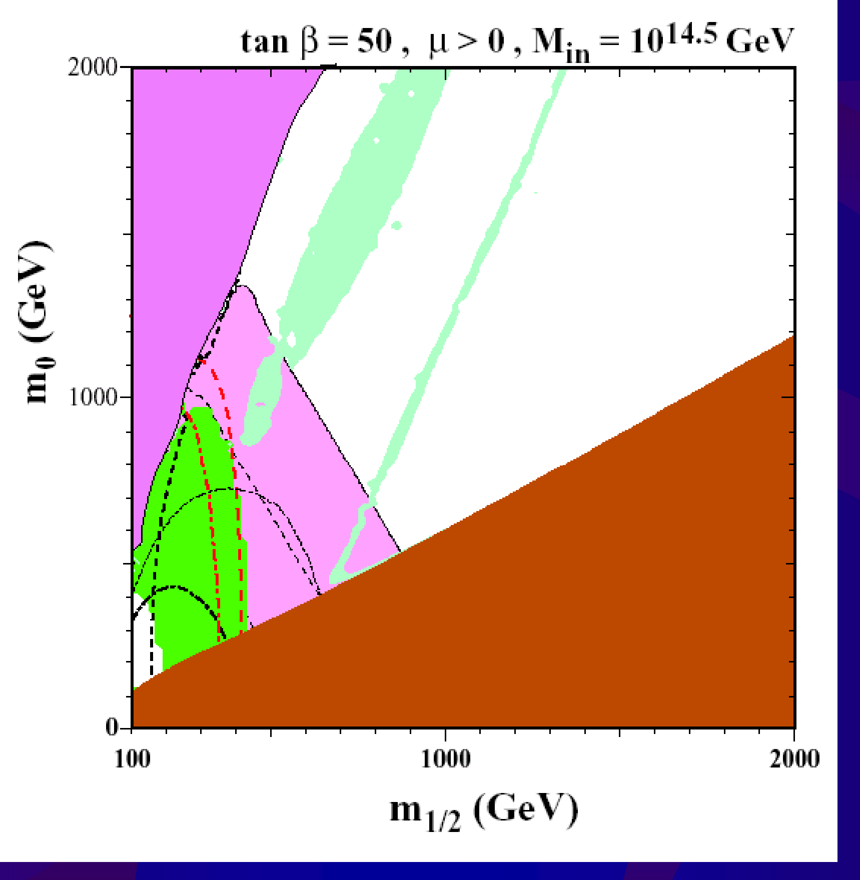
Lowering M_{in} - tan(β) = 50



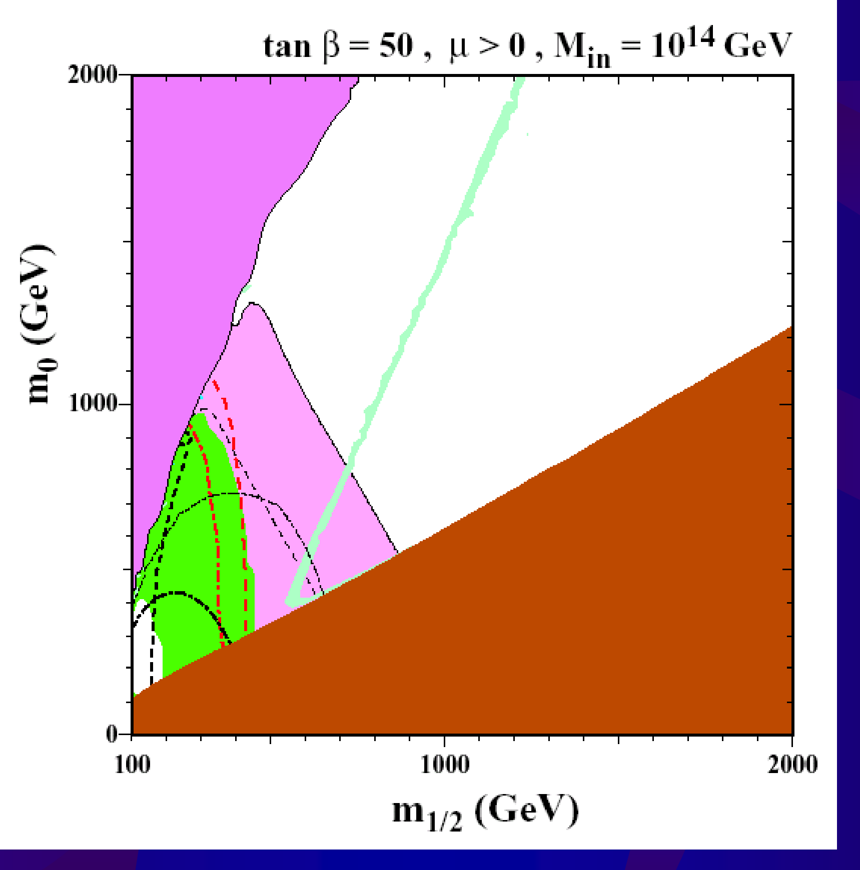
Lowering M_{in} - $\tan(\beta) = 50$



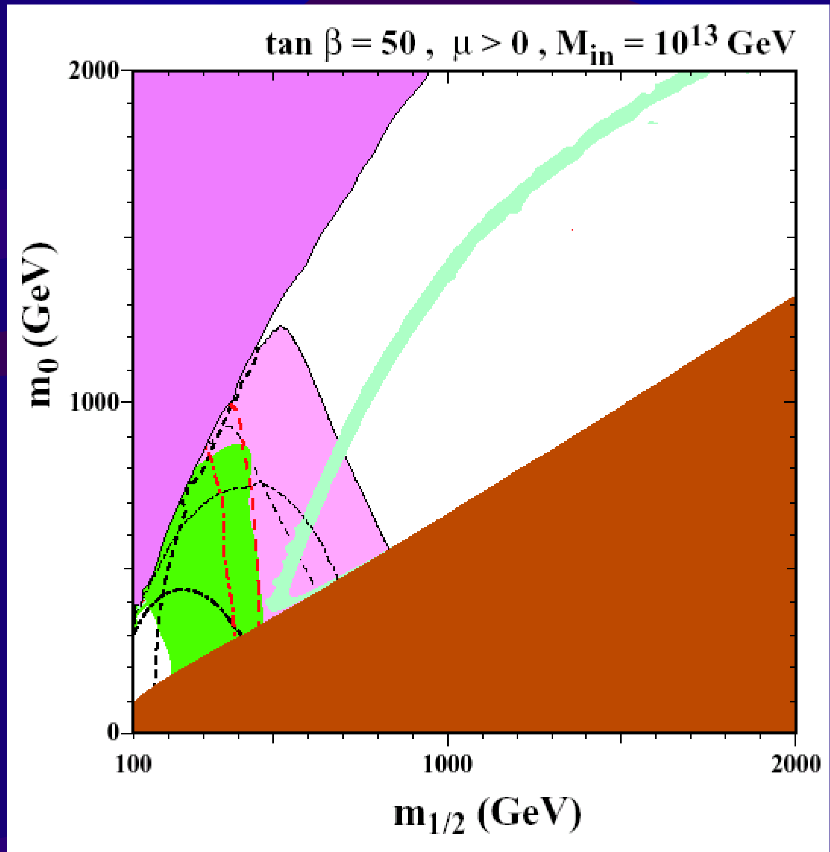
Lowering M_{in} - tan(β) = 50



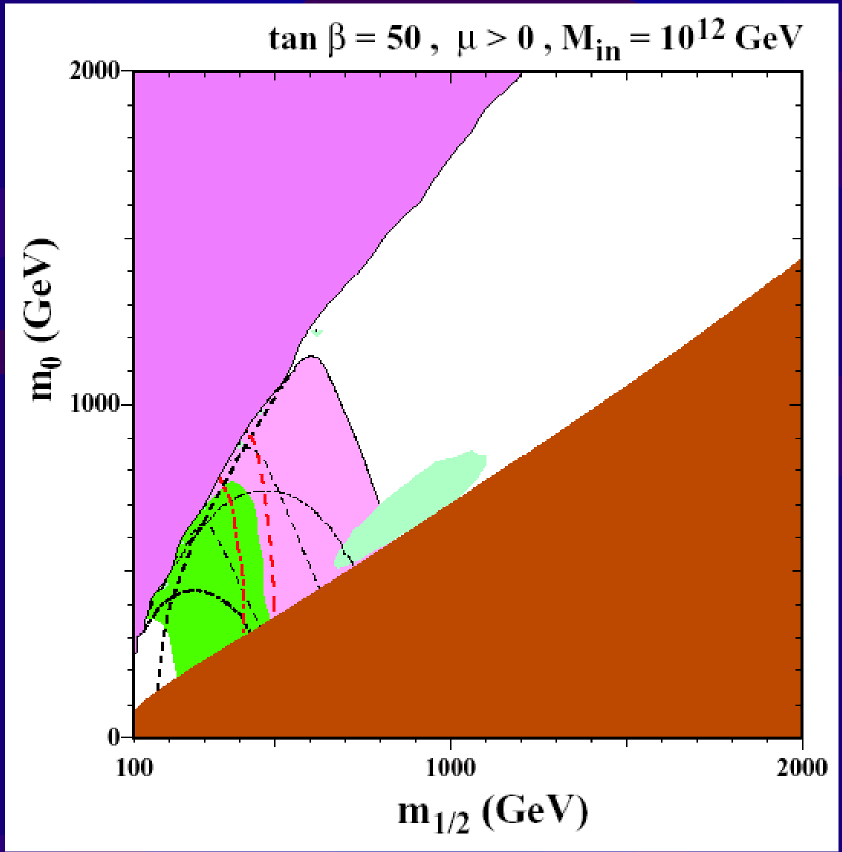
Lowering M_{in} - $\tan(\beta) = 50$



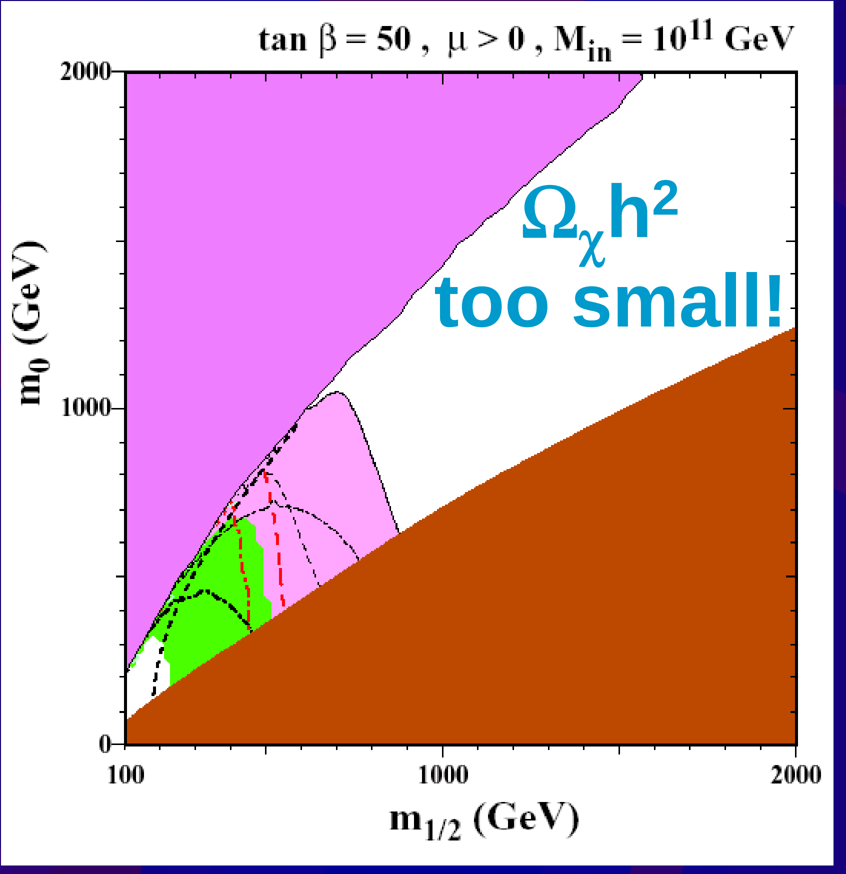
Lowering M_{in} - $\tan(\beta) = 50$



Lowering M_{in} - $\tan(\beta) = 50$

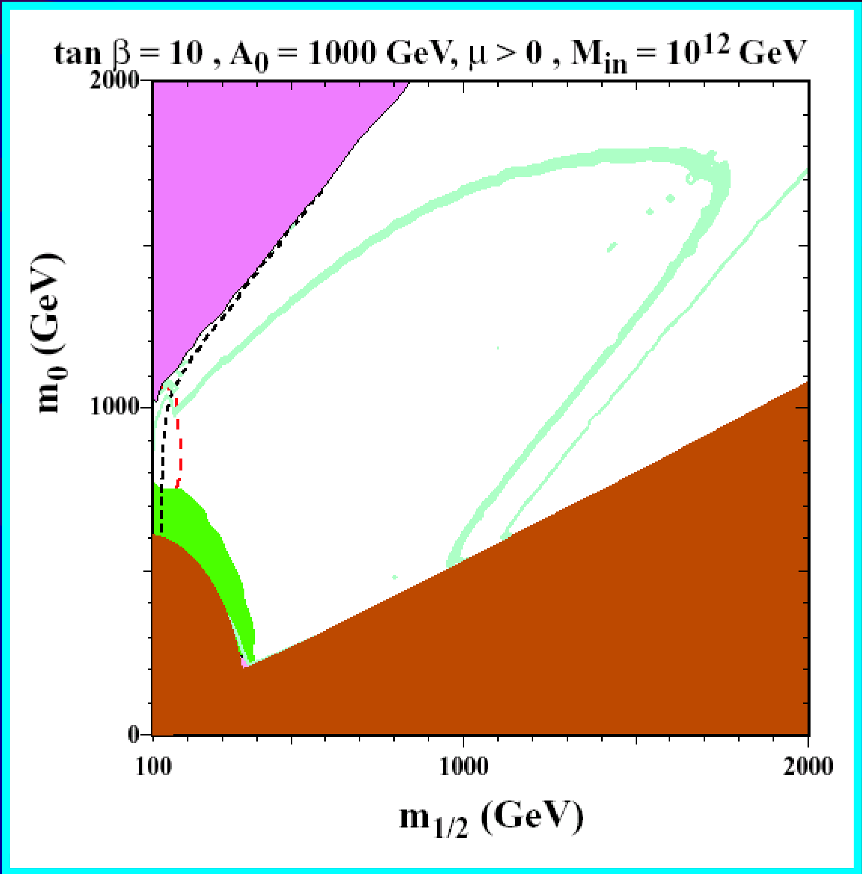


Lowering M_{in} - $\tan(\beta) = 50$

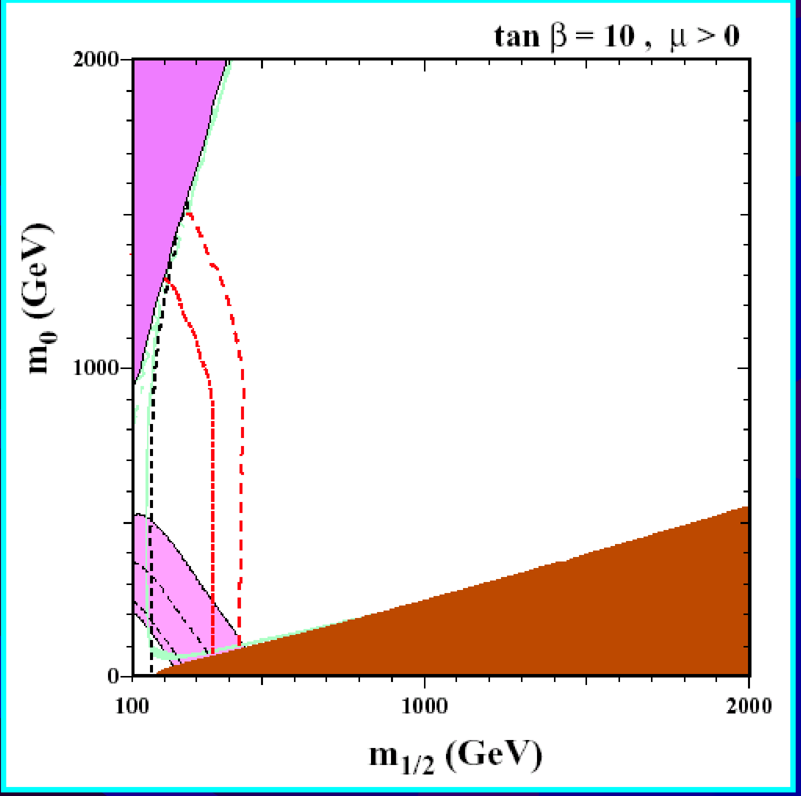


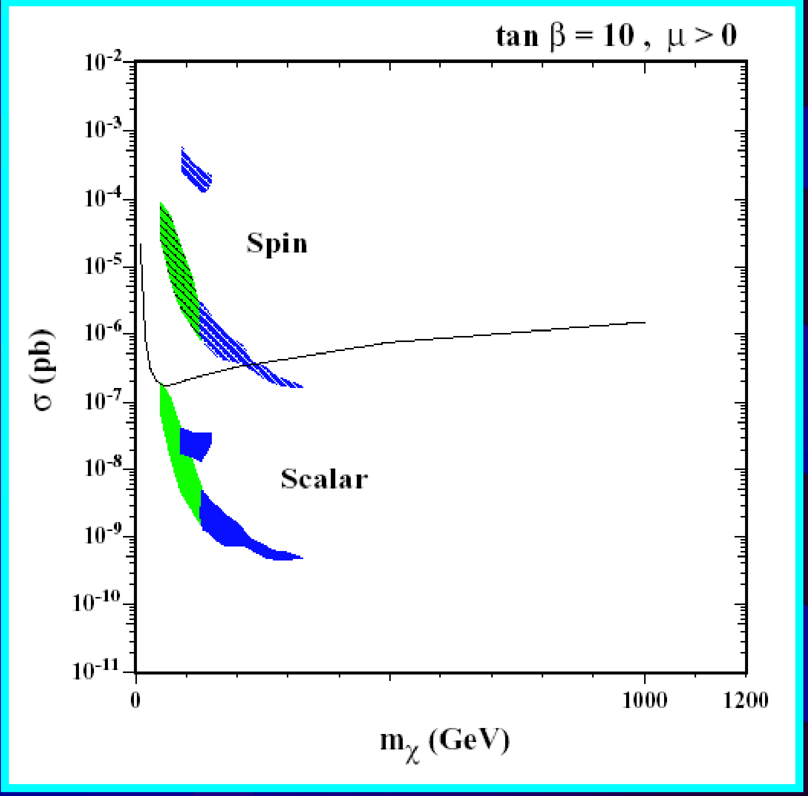
$A_0 \neq 0$

- $A_0 > 0 \Rightarrow$ larger weak-scale trilinear couplings, A_i
- Large loop corrections to μ depend on A_i, so μ is generically larger over the plane than when $A_0 = 0$.
- Also see stop-LSP excluded region

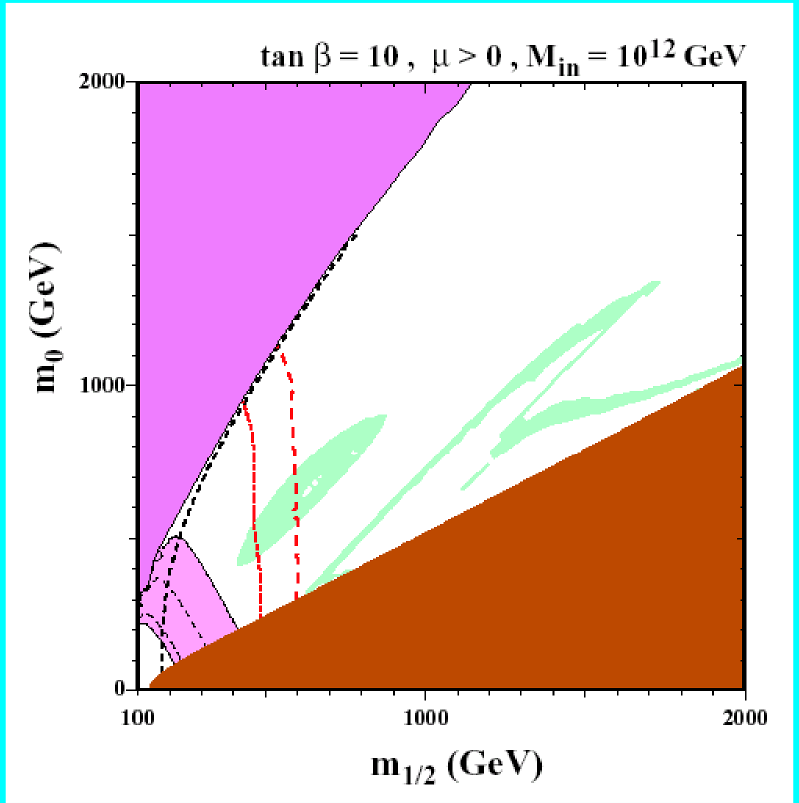


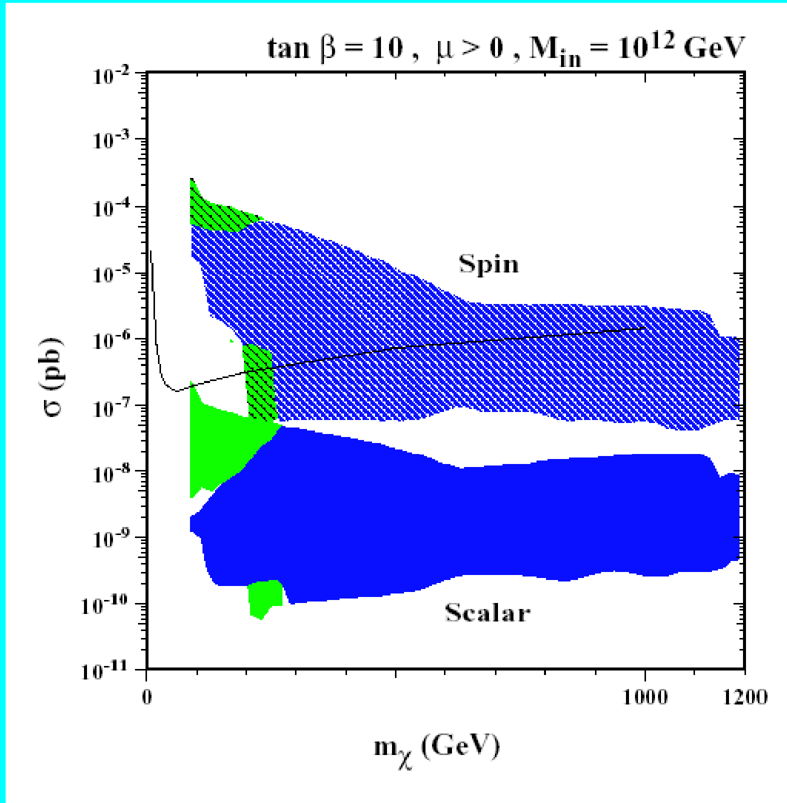
Direct Detection: Neutralino-Nucleon Cross Sections





Direct Detection: Neutralino-Nucleon Cross Sections

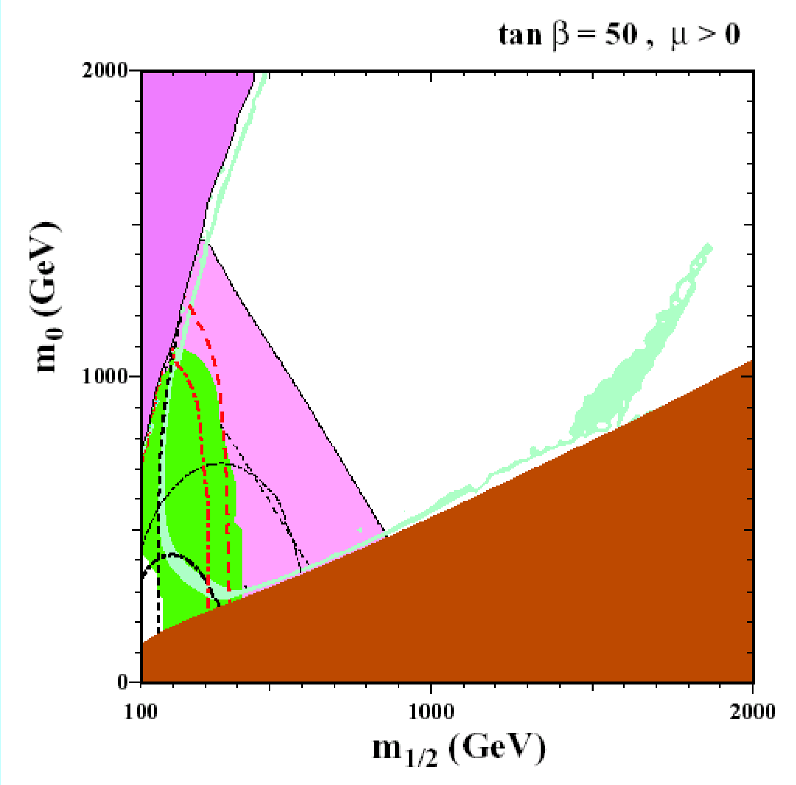


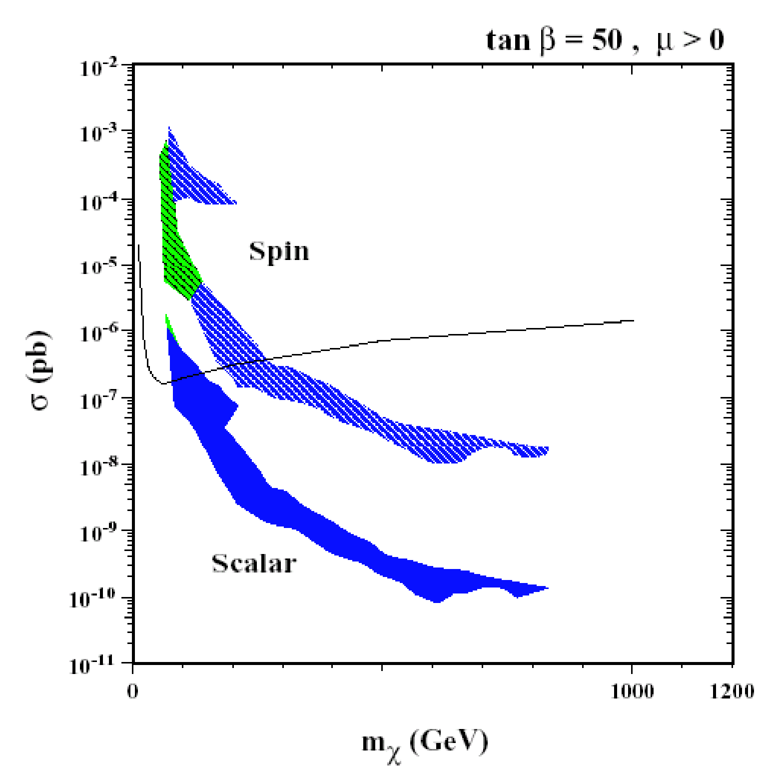


Conclusions

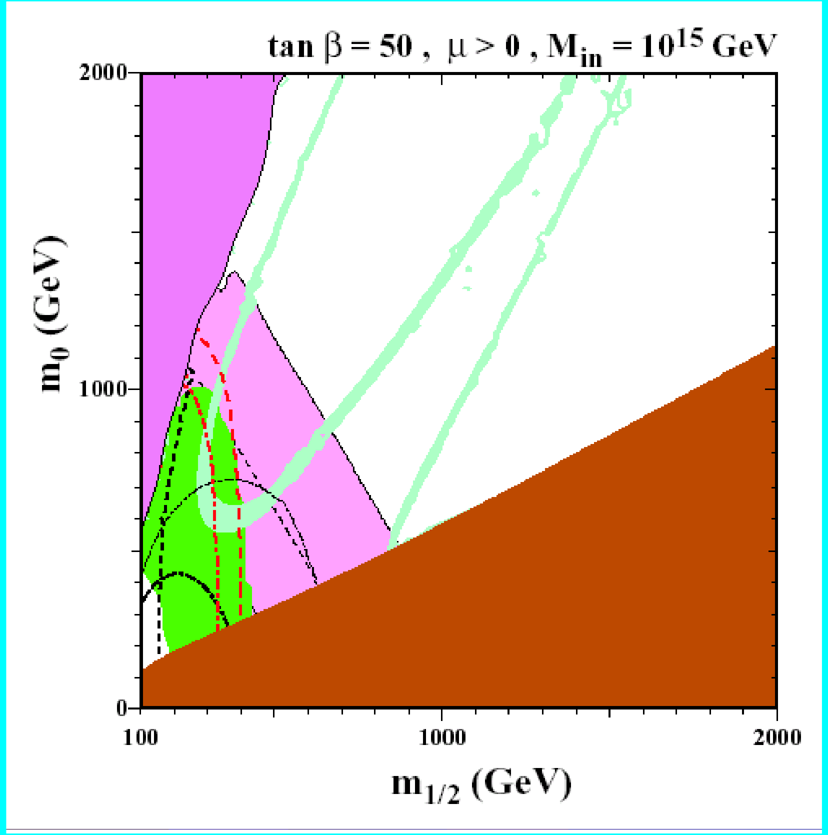
- Intermediate scale unification results in:
  - Rapid annihilation funnel even at low tan(β)
  - Merging of funnel and focus point
- Below some critical M_{in} (dependent on tan(β) and other factors), all or nearly all of the ($m_{1/2}$, m_0) plane is disfavored because the relic density of neutralinos is too low to fully account for the relic density of cold dark matter.

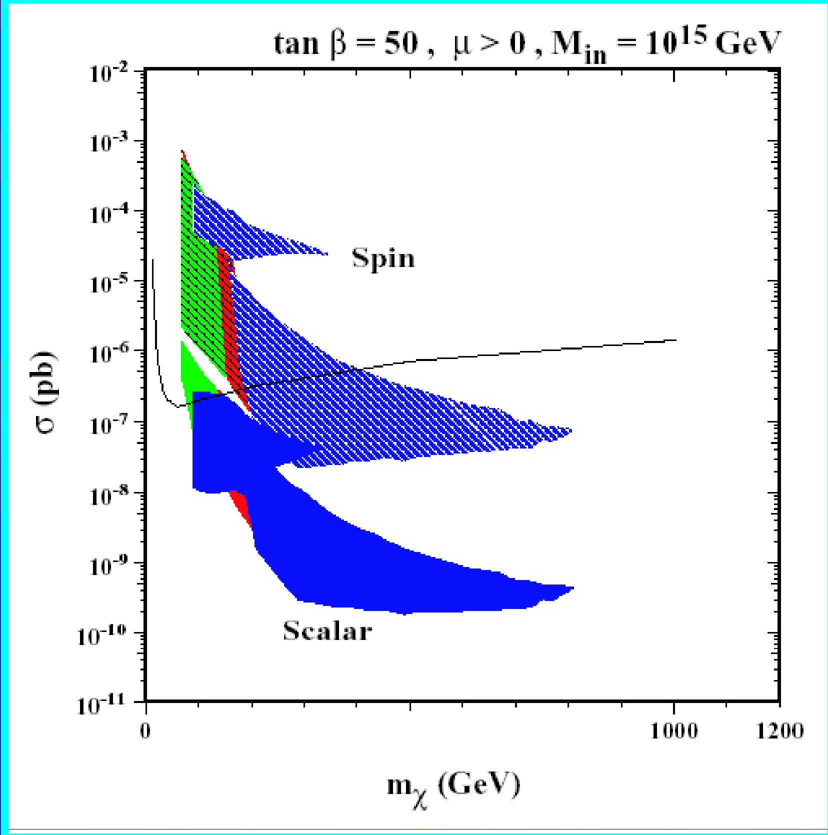
Neutralino-Nucleon Cross Sections





Neutralino-Nucleon Cross Sections





Sparticle Masses

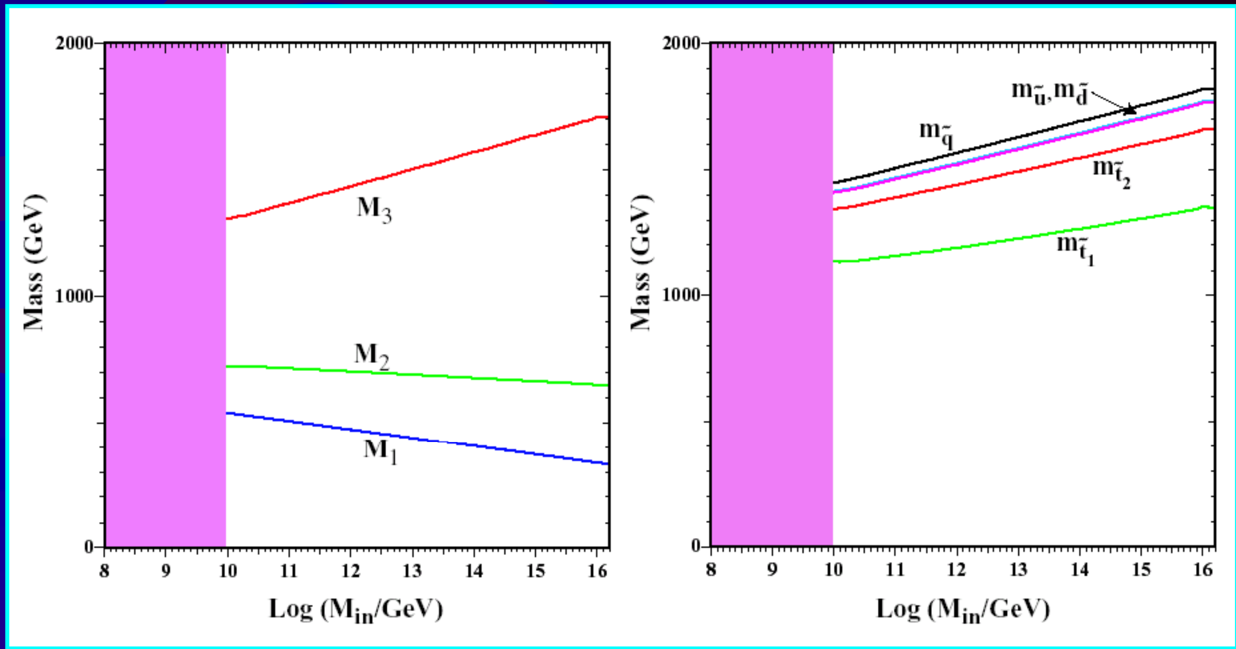
$m_{1/2}$ = 800 GeV

m_0= 1000 GeV

A_0 = 0

tan(β) = 10

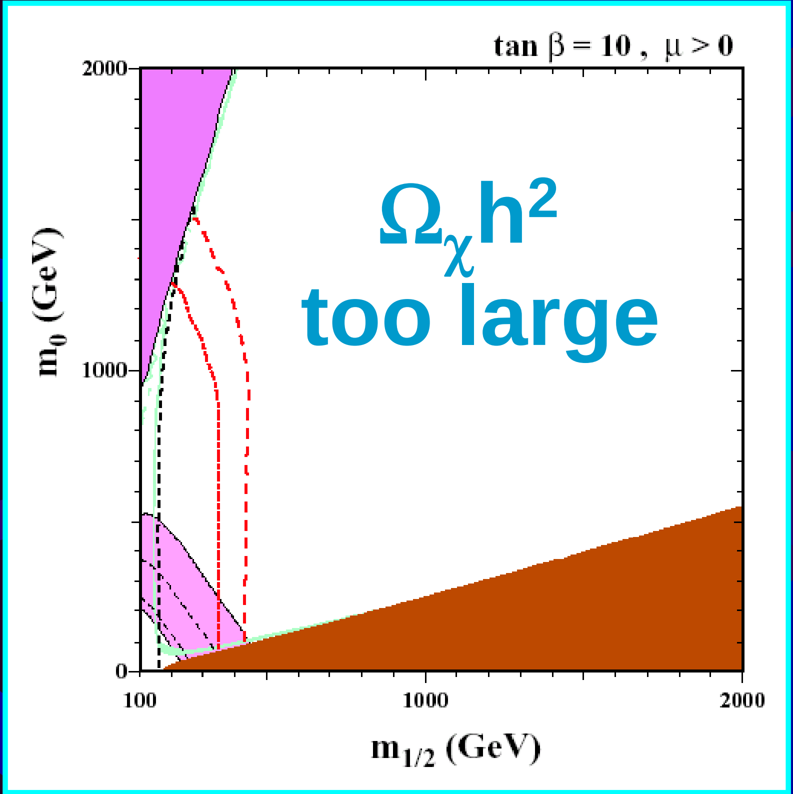
μ > 0

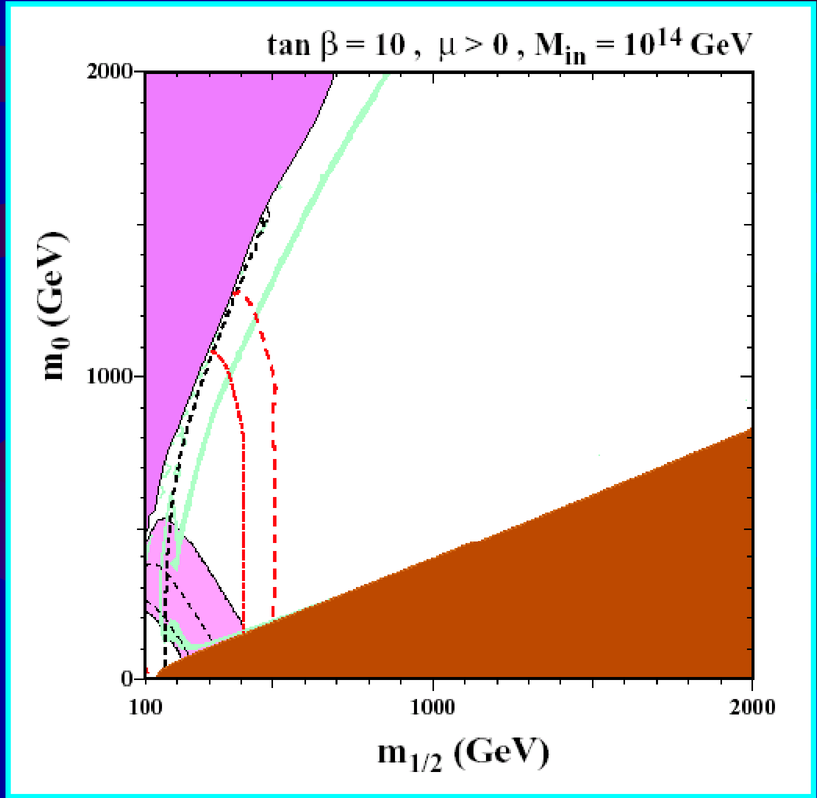


Guaginos

Squarks

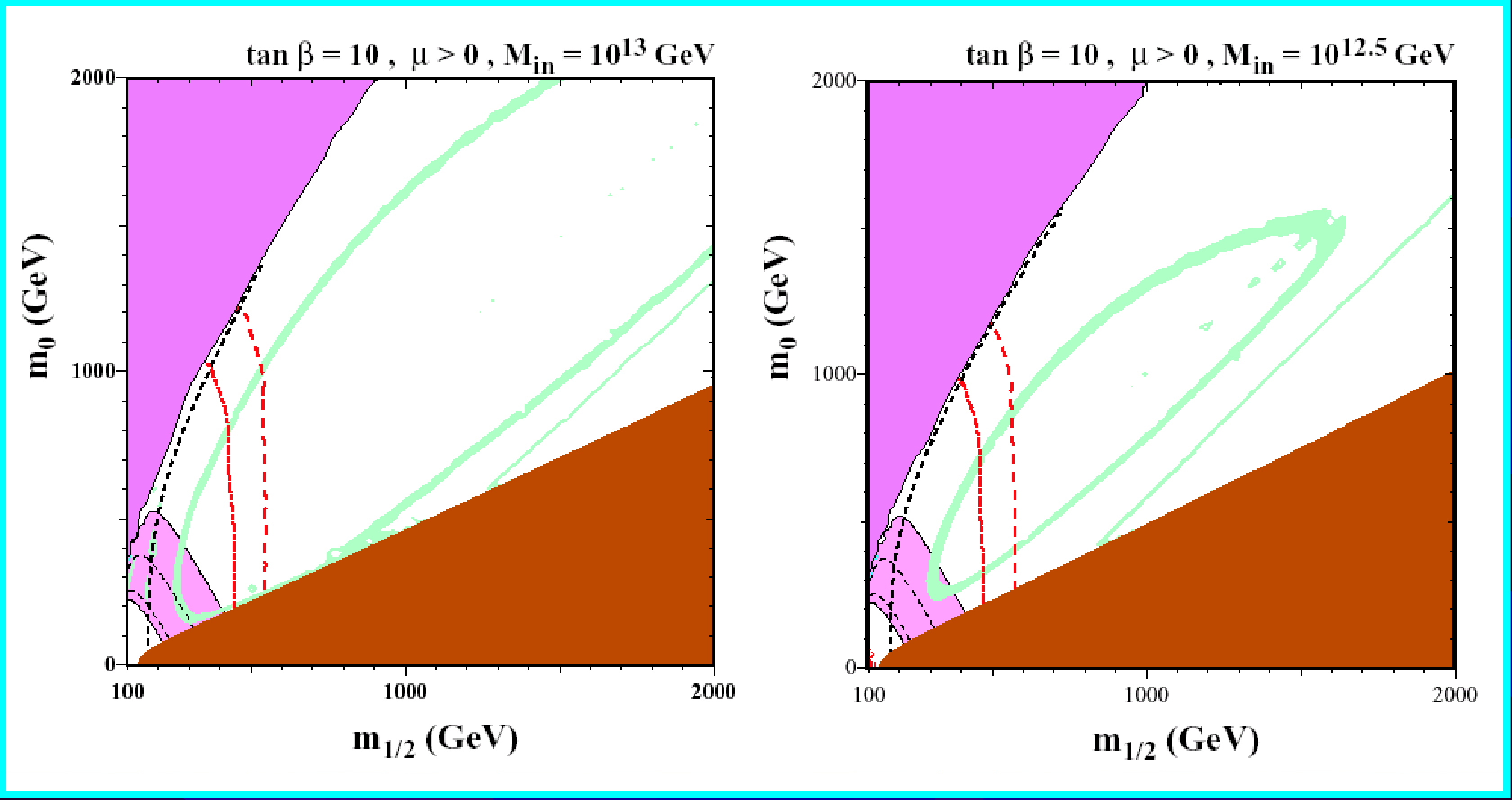
Lowering M_{in}



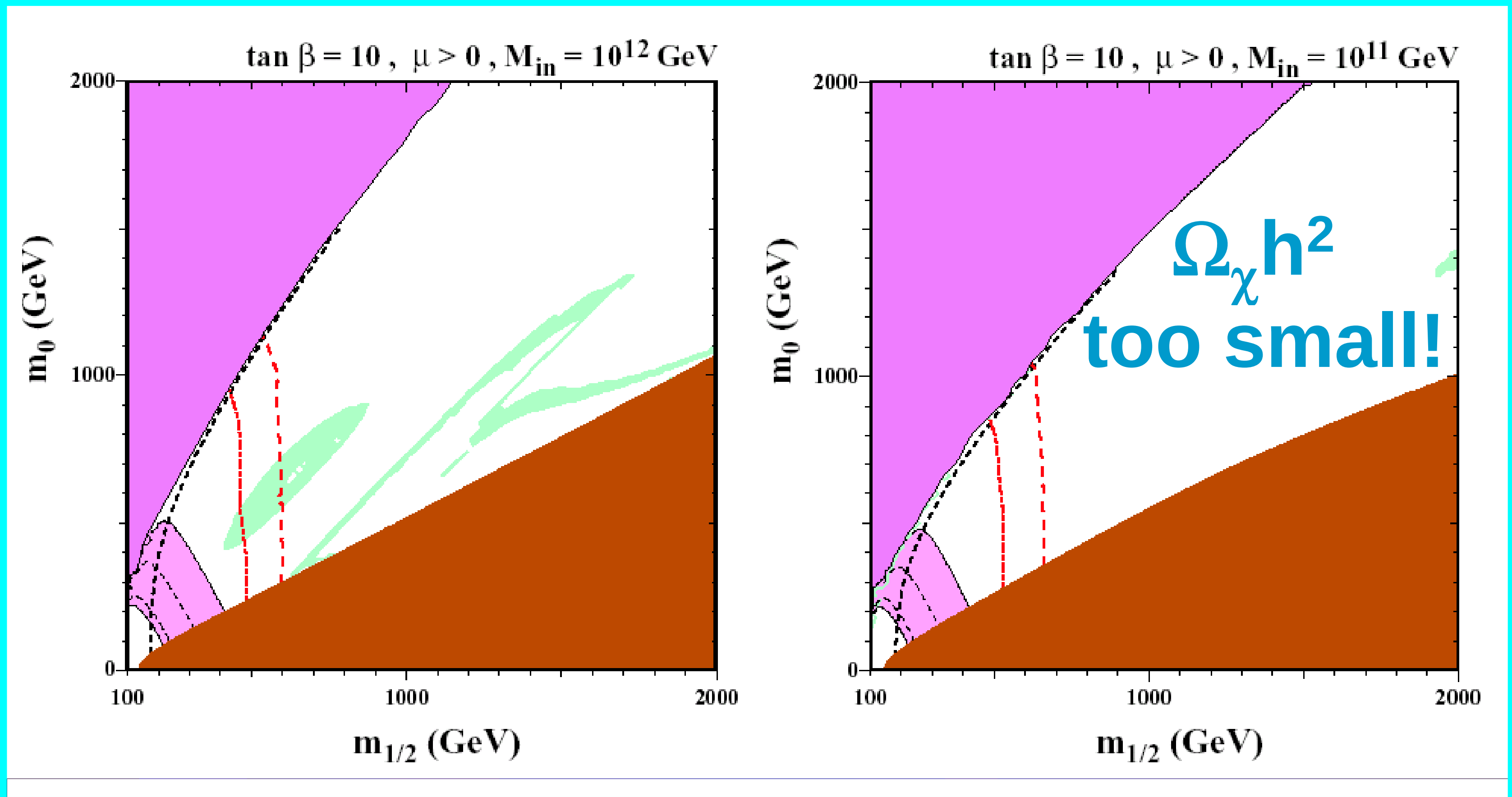


Pearl Sandick, UMN

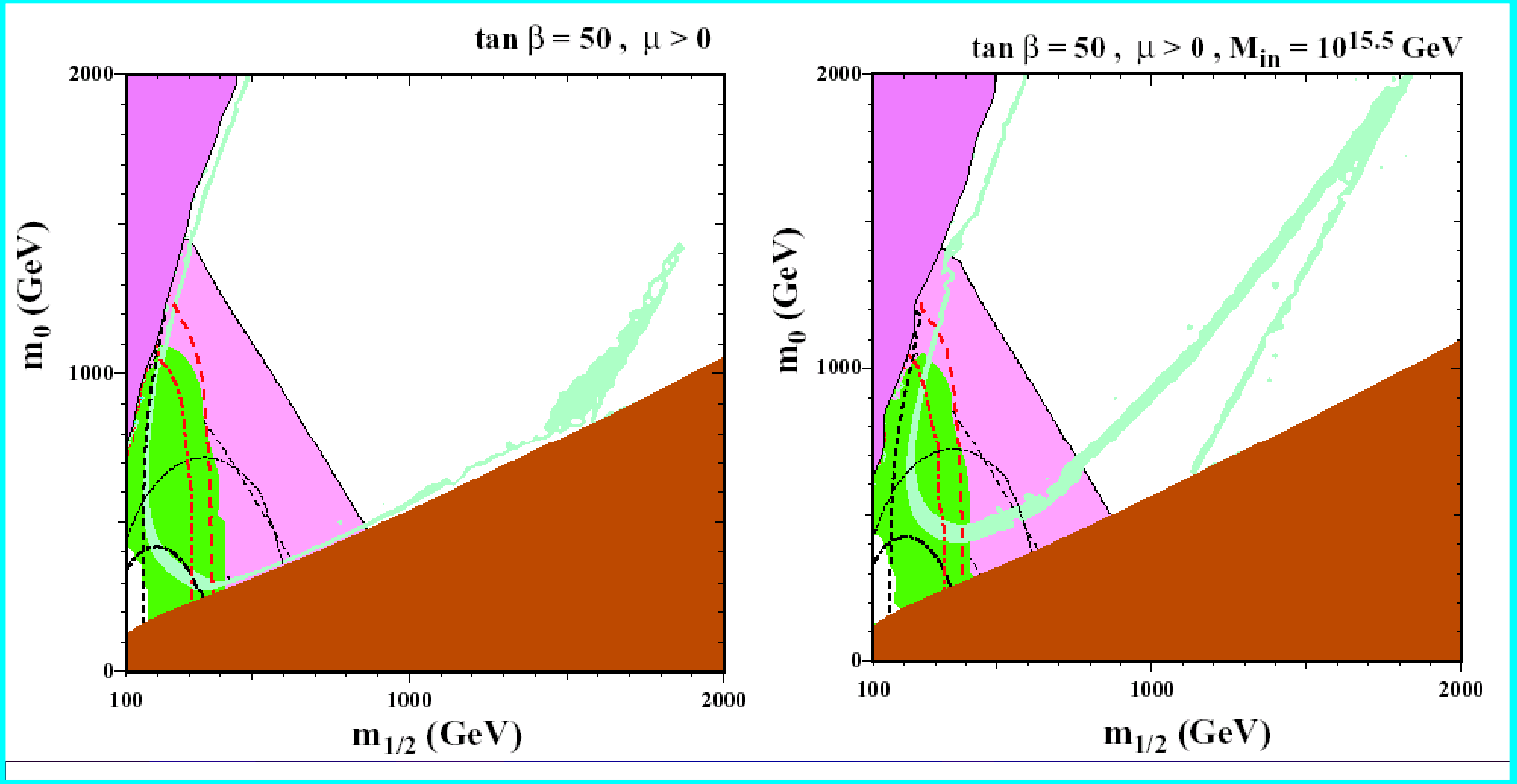
Lowering M_{in}



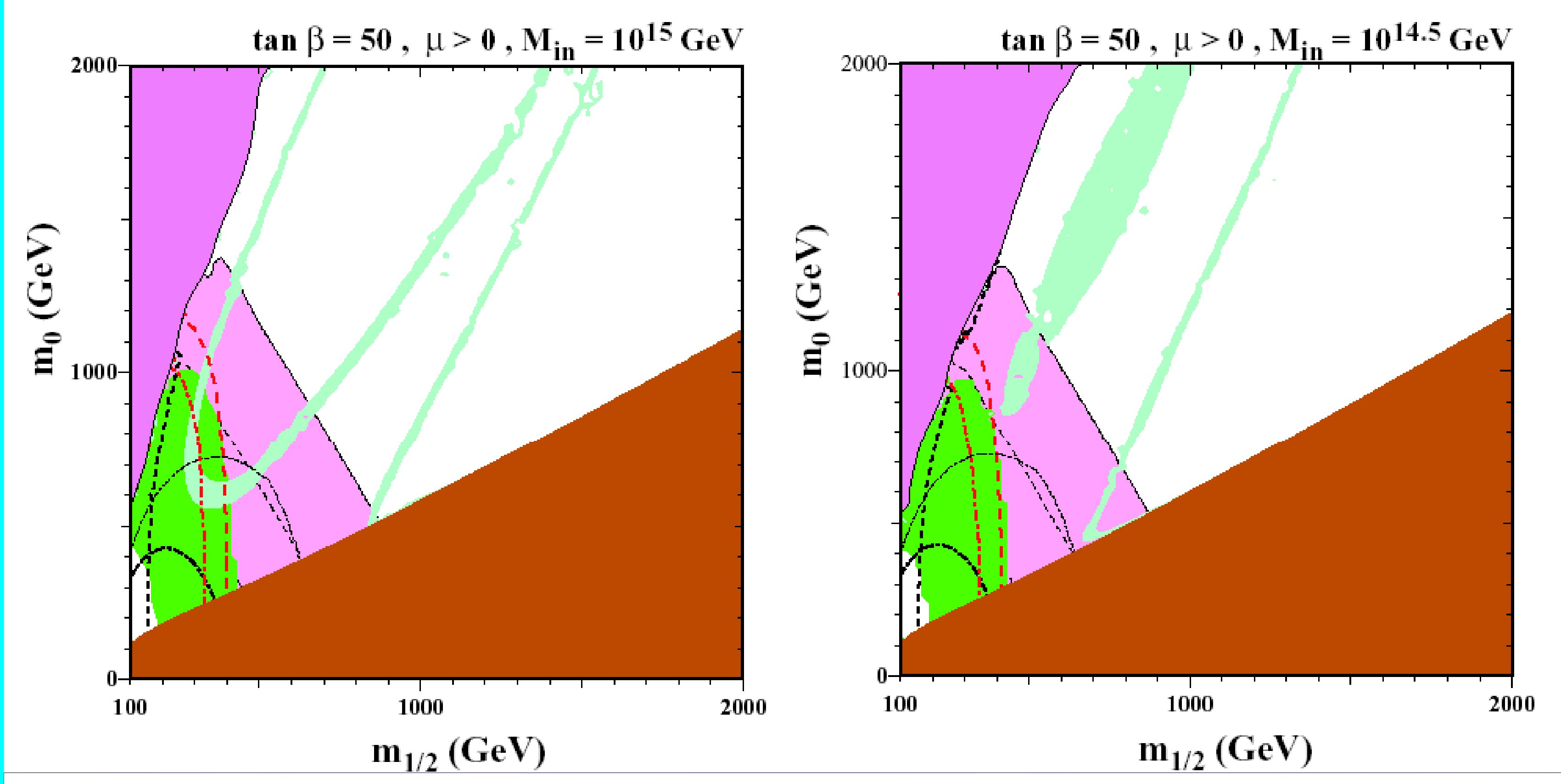
Lowering M_{in}



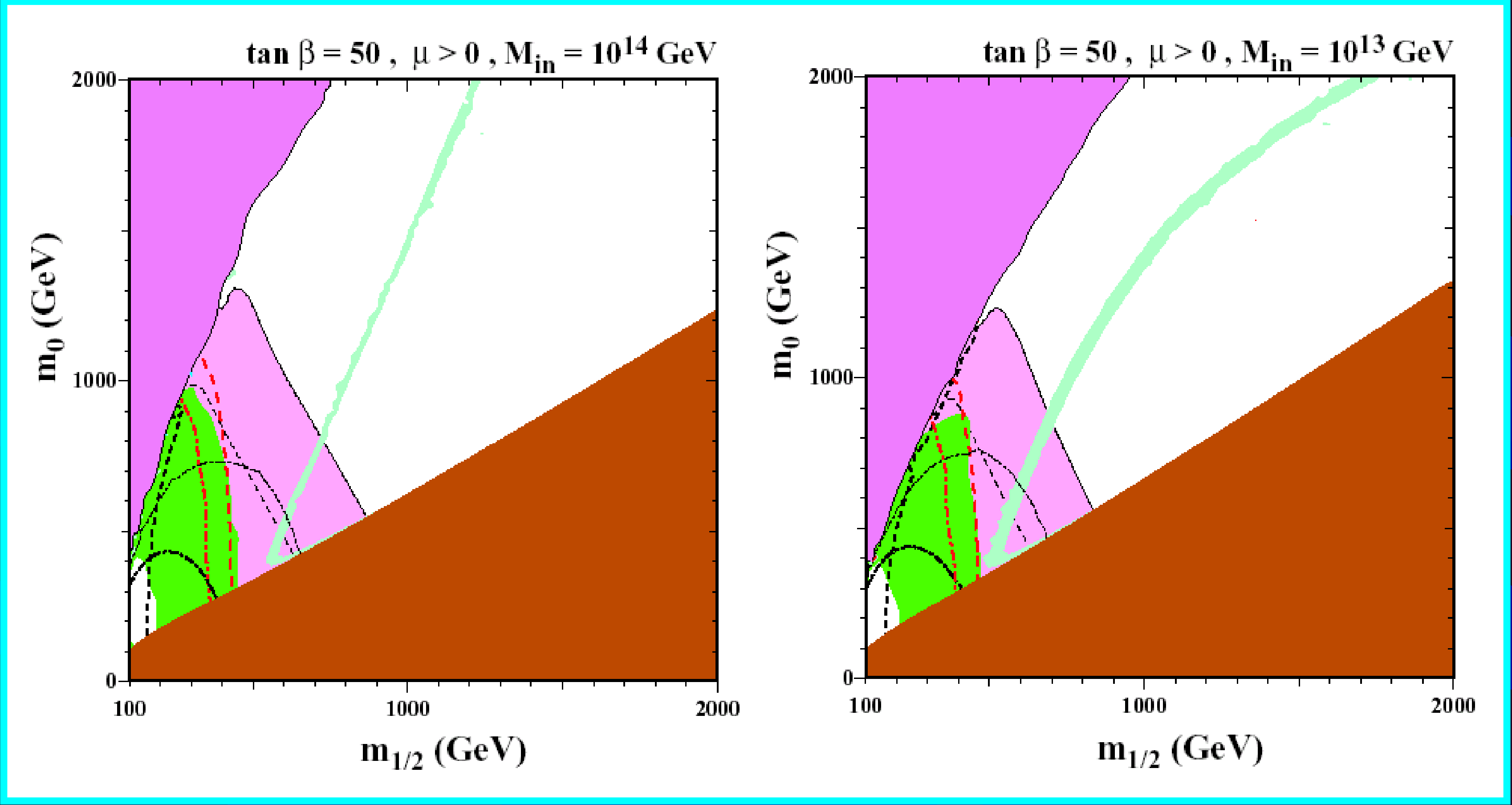
Lowering M_{in} - Large tan(β)



Lowering M_{in} - Large tan(β)



Lowering M_{in} - Large tan(β)



Lowering M_{in} - Large tan(β)