Radiative Inverse Seesaw and Dark Matter

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1. The Inverse Seesaw Mechanism
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Seesaw mechanism is one of the popular mechanism which can explain why neutrino masses are so much lighter than their charged lepton partners.

Where is the scale of seesaw, can it be close to the scale where LHC can say something? Usually larger than a few TeV

Inverse seesaw can further bring down the scale. If radiatively realized, the scale can be even lower. LHC may be able to say something about it.

Is neutrino mass generation mechanism related another fundamental problem of modern physics, dark matter? May be!

I will try convince you that indeed it is possibility to realize the above.

1. The Inverse Seesaw Mechanism

The Seesaw Mechanism

The Seesaw mechanism refers to the neutrino mass matrix of the form

$$L_m = -\frac{1}{2}(\nu_L^c, \nu_R)\begin{pmatrix} 0 & m_D \\ m_D^T & M_R \end{pmatrix}\begin{pmatrix} \nu_L \\ \nu_R^c \end{pmatrix}.$$

For one generation, if $M_R >> m_D$, the eigenmasses are

$$m_\nu \approx -m_D M_R^{-1} m_D^T, \quad m_N \approx M_R$$

A very nice way to explain why light neutrino masses are so much lighter than their charged lepton partners.

$$m_{\nu_1} = m_e^2/M_R \quad \rightarrow \quad M_R = m_e^2/m_{\nu_1}. \text{ For } m_{\nu_1} = 0.1\text{eV}, M_R = 2.5\text{TeV}$$
$$m_{\nu_2} = m_\mu^2/M_R \quad \rightarrow \quad M_R = m_\mu^2/m_{\nu_1}. \text{ For } m_{\nu_2} = 0.1\text{eV}, M_R = 10^8\text{GeV}$$
$$m_{\nu_3} = m_\tau^2/M_R \quad \rightarrow \quad M_R = m_\tau^2/m_{\nu_1}. \text{ For } m_{\nu_3} = 0.1\text{eV}, M_R = 3\times 10^{10}\text{GeV}.$$

Inverse Seesaw

The inverse seesaw neutrino mass matrix M_ν is the mass matrix resulted from the effective Lagrangian

$$L_m = -\bar{\nu}_L m_D N_R - \bar{N}_L M N_R - \frac{1}{2}\bar{N}^c_R \mu_R N_R - \frac{1}{2}\bar{N}_L \mu_L N^c_L + h.c.$$

where ν_L is the light active neutrino, $N_{L,R}$ are heavy neutrinos.

In the bases $(\nu^c_L, N_R, N^c_L)^T$, M_ν is given by

$$M_\nu = \begin{pmatrix} 0 & m_D & 0 \\ m_D^T & \mu_R & M^T \\ 0 & M & \mu_L \end{pmatrix}$$

With the hierarchy $\mu_L \sim \mu_R << m_D << M$, the light neutrino mass matrix $m_n u$, defined by $L_{mass} = -(1/2)\nu_L m_\nu \nu^c_L$, to order $(m_D/M)^2$ is given by

$$m_\nu = m_D M^{-1} \mu_L (M^{-1})^T m_D^T.$$

M can be much lower than that in usual Seesaw.

An example of Inverse Seesaw model

$$M_\nu = \begin{pmatrix} 0 & m_D & 0 \\ m_D^T & \mu_R & M^T \\ 0 & M & \mu_L \end{pmatrix}$$

introduce a leptonic doublet $D_{L,R} : (2, -1/2)$

to realize the inverse seesaw.

along with a singlet S and a triplet $\Delta : (3, -1)$

a global $U(1)_D$ symmetry to distinguish D_L and L_L.

Under this symmetry $D_{L,R} \to exp[i\alpha_D]D_{L,R}$, $S \to exp[-i\alpha_D]S$, $\Delta \to exp[2i\alpha_D]\Delta$,

and other fields do not transform.

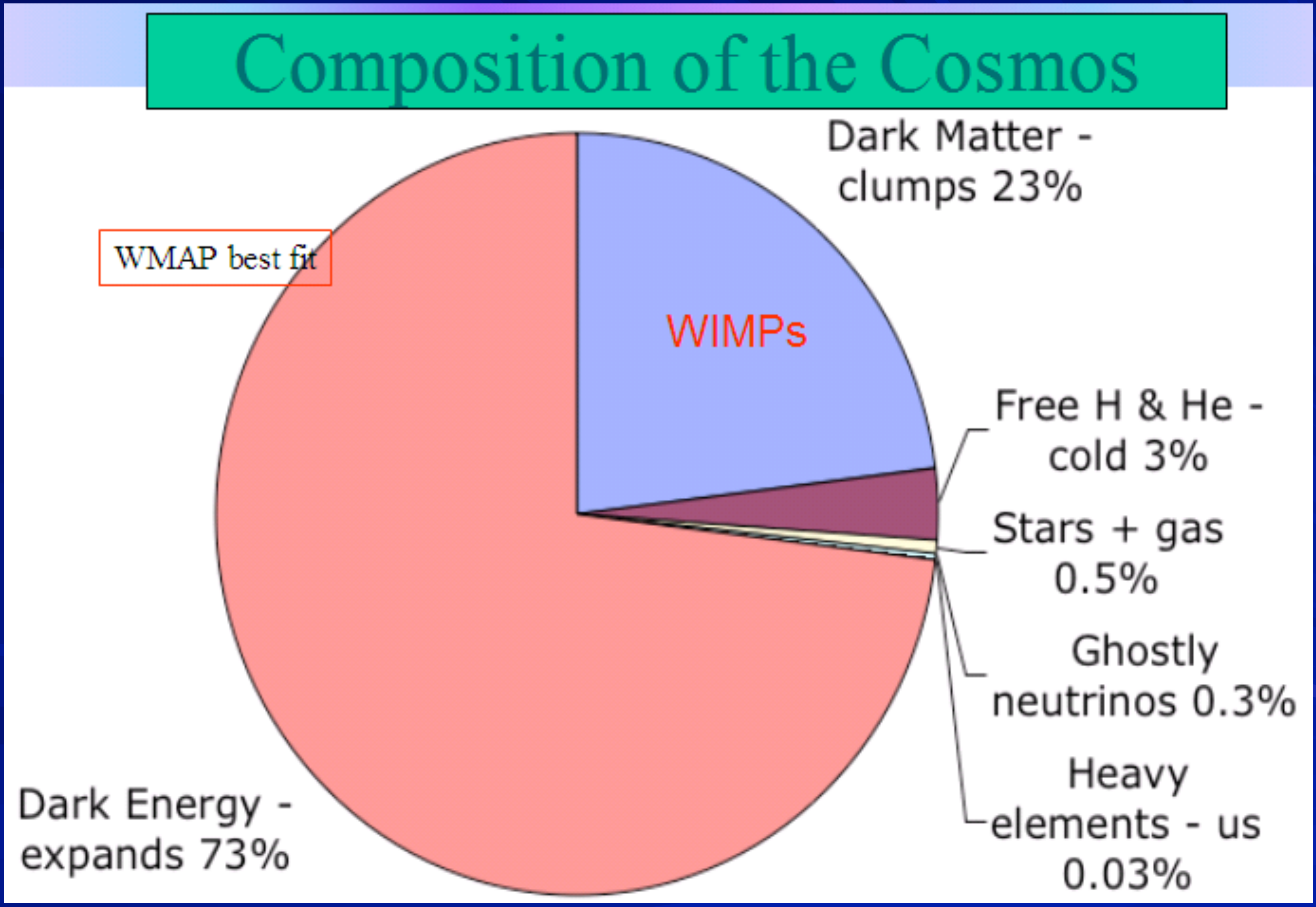
$$L_D = -\bar{L}_L Y_D D_R S - \bar{D}_L M D_R - \frac{1}{2}\bar{D}_L Y_L D_L^c \Delta - \frac{1}{2}\bar{D}_R^c Y_L D_R \Delta^\dagger + h.c.$$

If both S and Δ both develop non-zero vev's, the inverse seesaw mechanism is realized. This model, however, will have a Goldstone boson due to breaking of the global $U(1)_D$ symmetry which may be problematic.

To avoid the existence of a Goldstone boson in the theory, extension is needed.

Introduce more fields, or softly break $U(1)_D$ by adding a term of the form HDH

Link Neutrino Physics to Dark Matter physics



Usual seesaw, N (heavy neutrino) mixes with light ν_L, although the mixing is small of order sqrt[m_ν/m_N], but it is still too large to keep N long lived to be candidate for dark matter (smaller mixing may do the job with small N mass), N cannot be WIMP dark matter candidate.

Need additional symmetry to guarantee the stability of N.

Radiative seesaw and dark matter

E. Ma, PRD73:077301,2006 introduce right handed neutrinos N_R and a new Higgs doublet h

$$SU(2)_L \times U(1)_Y \bar{\times} Z_2$$

$$(\nu_i, l_i) \sim (2, -1/2; +), \quad l_i^c \sim (1, 1; +), \quad N_i \sim (1, 0; -),$$

$$(\phi^+, \phi^0) \sim (2, 1/2; +), \quad (\eta^+, \eta^0) \sim (2, 1/2; -).$$

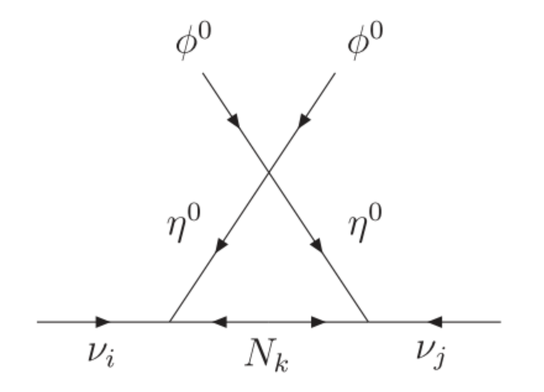
$$\mathcal{L}_Y = f_{ij}(\phi^- \nu_i + \bar{\phi}^0 l_i) l_j^c + h_{ij}(\nu_i \eta^0 - l_j \eta^+) N_j + \text{H.c.}$$

$$\tfrac{1}{2} M_i N_i N_i + \text{H.c.} \qquad \tfrac{1}{2}\lambda_5 (\Phi^\dagger \eta)^2 + \text{H.c.}$$

under Z_2

N_R -> - N_R

η -> - η



$$(\mathcal{M}_\nu)_{ij} = \sum_k \frac{h_{ik} h_{jk} M_k}{16\pi^2} \left[\frac{m_R^2}{m_R^2 - M_k^2} \ln \frac{m_R^2}{M_k^2} - \frac{m_I^2}{m_I^2 - M_k^2} \ln \frac{m_I^2}{M_k^2} \right],$$

If $m_0^2 \simeq M_k^2$, then $\quad (\mathcal{M}_\nu)_{ij} \simeq \frac{\lambda_5 v^2}{16\pi^2} \sum_k \frac{h_{ik} h_{jk}}{M_k}.$

η no vev, Z_2 not broken. No usual seesaw neutrino masses. One loop will generate them.

If neutral h is the lightest particle, it can be DM

If one of the N is the lightest, it is the DM

DM mass related to radiative seesaw neutrino mass scale.

Radiative Seesaw can further lower the seesaw scale and having dark matter candidate as a bonus!

What about Radiative inverse seesaw with dark matter?

2. Two Loop Radiative Inverse Seesaw

Besides the problem of Goldstone boson, the example model given earlier does not have dark matter candidates.

To avoid the existence of Goldstone bosons, do not let relevant scalars to have non-zero vevs (without adding soft breaking terms). It implies that one cannot have vev for S and Δ. The global U(1) symmetry is not broken.

Bonus: Unbroken symmetry -> guarantee
the existence of stable particles -> dark matter candidate.

Problem: light neutrinos remains massless.
Solution: Radiatively generate neutrino masses.

The U(1)D model

introduce a leptonic doublet $D_{L,R} : (2, -1/2)$

to realize the inverse seesaw.

along with a singlet S and a triplet $\Delta : (3, -1)$

a global $U(1)_D$ symmetry to distinguish D_L and L_L.

Under this symmetry $D_{L,R} \to exp[i\alpha_D]D_{L,R}$, $S \to exp[-i\alpha_D]S$, $\Delta \to exp[2i\alpha_D]\Delta$,

and other fields do not transform.

$$L_D = -\bar{L}_L Y_D D_R S - \bar{D}_L M D_R - \frac{1}{2}\bar{D}_L Y_L D_L^c \Delta - \frac{1}{2}\bar{D}_R^c Y_L D_R \Delta^\dagger + h.c.$$

If both S and Δ both develop non-zero vev's, the inverse seesaw mechanism is realized. This model, however, will have a Goldstone boson due to breaking of the global $U(1)_D$ symmetry which may be problematic.

To avoid the existence of a Goldstone boson in the theory, extension is needed.

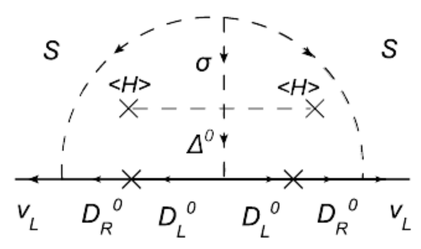
Cure:

introduce another singlet σ which transforms under the $U(1)_D$ as $\sigma \rightarrow exp[2i\alpha_D]\sigma$.

The allowed renormalizable terms in the potential V_D are given by

$$V_D = -\mu_H^2 H^\dagger H + \lambda_H (H^\dagger H)^2 + \mu_S^2 S^\dagger S + \lambda_S (S^\dagger S)^2 + \mu_\sigma^2 \sigma^\dagger \sigma + \lambda_\sigma (\sigma^\dagger \sigma)^2$$
$$+ \mu_\Delta^2 \Delta^\dagger \Delta + \lambda_\Delta^\alpha (\Delta^\dagger \Delta \Delta^\dagger \Delta)_\alpha + \sum_{ij} \lambda_{ij} i^\dagger i j^\dagger j + (\mu_{S\sigma} S^2 \sigma + \lambda_{\Delta\sigma H} H \Delta \sigma^\dagger H + h.c.)$$

where the sum $\sum_{ij}$ is over all possible i and j to be one of the H, S, σ and Δ.



The two loop diagram break lepton number, but conserve U(1)D number

In the above μ_i^2 are all larger than zero. The potential only allows H to have a non-zero vev v_H. The theory after spontaneous symmetry breaking from $SU(2)_L \times U(1)_Y$ to $U(1)_{em}$, the $U(1)_D$ global symmetry is unbroken.

At the tree level, light neutrinos are massless.

Non-zero neutrino masses can only be generated at two loop level

This is similar to the Babu-Zee model two loop neutrino mass generation

but with the light charged leptons in the loop replaced by new heavy particles.

in the bases where M is diagonalized
$$m_\nu^{ij} = \frac{v_H Y_D^{ik} (\lambda_{\Delta\sigma H} \mu_{S\sigma} Y_L^{kl}) Y_D^{jl} v_H}{M_{kk}^2} \kappa_{kl} \ ,$$

κ_{kl} is defined as:
$$\kappa_{kl} = \delta_{kl} \frac{1}{2(4\pi)^4} \frac{1}{(1 - m_S^2/M_{kk}^2)^2} [g(m_{\phi_1}, m_S, m_S) - g(m_{\phi_1}, M_{kk}, m_S) \\ -g(m_{\phi_1}, m_S, M_{kk}) + g(m_{\phi_1}, M_{kk}, M_{kk})]$$

$$g(m_1, m_2, m_3) = \int_0^1 dx [1 + Sp(1 - \mu^2) - \frac{\mu^2}{1 - \mu^2} \log \mu^2] \qquad Sp(z) = - \int^z \frac{\ln(1-t)}{t} dt$$

with $\mu^2 = \frac{ax + b(1-x)}{x(1-x)}, a = \frac{m_2^2}{m_1^2}, b = \frac{m_3^2}{m_1^2}$.

$$M_\nu = \begin{pmatrix} 0 & m_D & 0 \\ m_D^T & \mu_R & M^T \\ 0 & M & \mu_L \end{pmatrix}$$

If one identify, effectively,

$m_D = Y_D v_H$, $M = diag(M_{ii})$ and $\mu_L = (\mu_L^{ij})$ $\mu_L^{ij} = (\lambda_{\Delta\sigma H} \mu_{s\sigma}) Y_L^{ij} \kappa_{ij}$,

the light neutrino mass matrix is effectively an inverse seesaw mass form.

We therefore refer this as radiative inverse seesaw mechanism.

Fitting to data

Neutrino mixing and mass differences from Global data fitting

Fogli et al, arXiv: 1205:5204

Results of the global 3ν oscillation analysis, in terms of best-fit values and allowed 1, 2 and 3σ ranges for the 3ν mass-mixing parameters. We remind that Δm^2 is defined herein as $m_3^2 - (m_1^2 + m_2^2)/2$, with $+\Delta m^2$ for NH and $-\Delta m^2$ for IH.

Parameter	Best fit	1σ range	2σ range	3σ range
$\delta m^2/10^{-5}$ eV^2 (NH or IH)	7.54	7.32 – 7.80	7.15 – 8.00	6.99 – 8.18
$\sin^2\theta_{12}/10^{-1}$ (NH or IH)	3.07	2.91 – 3.25	2.75 – 3.42	2.59 – 3.59
$\Delta m^2/10^{-3}$ eV^2 (NH)	2.43	2.33 – 2.49	2.27 – 2.55	2.19 – 2.62
$\Delta m^2/10^{-3}$ eV^2 (IH)	2.42	2.31 – 2.49	2.26 – 2.53	2.17 – 2.61
$\sin^2\theta_{13}/10^{-2}$ (NH)	2.41	2.16 – 2.66	1.93 – 2.90	1.69 – 3.13
$\sin^2\theta_{13}/10^{-2}$ (IH)	2.44	2.19 – 2.67	1.94 – 2.91	1.71 – 3.15
$\sin^2\theta_{23}/10^{-1}$ (NH)	3.86	3.65 – 4.10	3.48 – 4.48	3.31 – 6.37
$\sin^2\theta_{23}/10^{-1}$ (IH)	3.92	3.70 – 4.31	3.53 – 4.84 $\oplus$ 5.43 – 6.41	3.35 – 6.63
δ/π (NH)	1.08	0.77 – 1.36	—	—
δ/π (IH)	1.09	0.83 – 1.47	—	—

Example for neutrino mss solutions

For the normal hierarchy, choosing $\hat{Y}_D = diag(1, \sqrt{1.03}, \sqrt{1.77})$, $y_D \times \lambda_{\Delta\sigma H} = 10^{-3}$, $Y_L = \mathrm{I} \times 10^{-2}$, $\mu_{S\sigma} = 100\mathrm{GeV}$, $m_{\phi_1} = 300\mathrm{GeV}$, $m_S = 150\mathrm{GeV}$, $M_{ii} = 500\mathrm{GeV}$, we can get all the three neutrino mass $3.39 \times 10^{-2}\mathrm{eV}$, $3.50 \times 10^{-2}\mathrm{eV}$, $5.98 \times 10^{-2}\mathrm{eV}$, respectively. These are consistent with data.

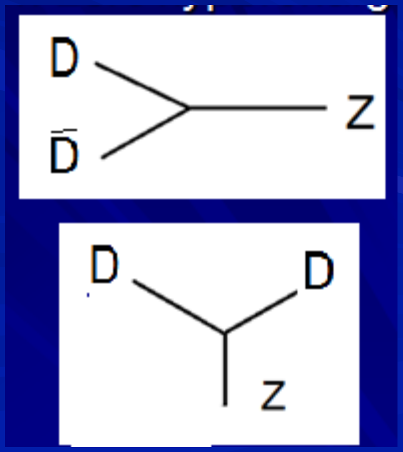
For inverted hierarchy case, we just need to replace $\hat{Y}_D$ with $\hat{Y}_D = diag(\sqrt{1.46}, \sqrt{1.48}, \sqrt{0.100})$, with all the other parameters unchanged, the neutrino masses will be $4.93 \times 10^{-2}\mathrm{eV}$, $5.01 \times 10^{-2}\mathrm{eV}$, $3.39 \times 10^{-3}\mathrm{eV}$, respectively. Again, these numbers are consistent with data.

3. Dark Matter in Radiative Inverse Seesaw

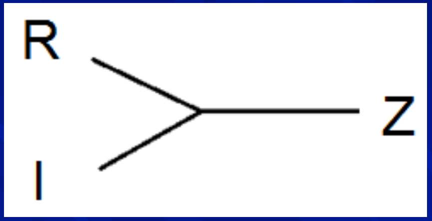
In the $U(1)_D$ models, since the global symmetry $U(1)_D$ is not broken, there is at least one stable new particle in the model. They are potential dark matter candidates.

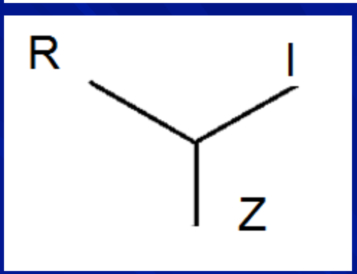
The $U(1)_D$ model: S, σ, and neutral component in D and Δ.

The neutral component in D has hypercharge and has a non-zero coupling to Z. Can fit relic density through s-channel Z exchange, but leads to a too large direct detection cross section through t-channel Z exchange. Out!



Neutral components in Δ also have non-zero hypercharge and couple to Z-Re(Δ)-Im(Δ).





If there is a mass splitting $\delta = m_R - m_I$ larger than 100 KeV or so between Re(Δ) and Im(Δ), It is possible to satisfy relic density and direct detection constraints, through inelastic dark matter idea.

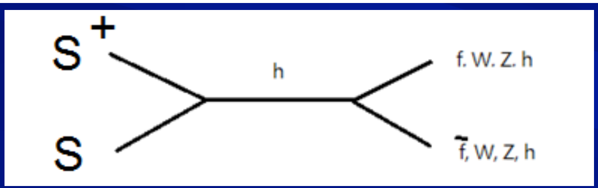
But in both models, there is no mass splitting, $\delta = 0$, inelastic dark matter mechanism is ineffective. Out!

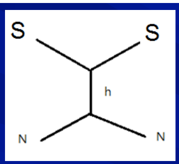
σ mixes with neutral component in Δ, if mixing is large, ruled out. Also, there is a Z_2 unbroken symmetry for S -> -S, Δ -> -Δ and others do not transform, S or neutral Δ is also stable. Two component dark matter, σ and S. (Δ is out due to non-zero hypercharge)

S in U(1)D model can be a natural dark matter candidate. It is a complex singlet dark matter model. It is a Higgs portal dark matter model. The relic density and direct detection is through s- and t-channel higgs h exchange, respectively, via,

$$\lambda_{SH} S^\dagger S H^\dagger H = \frac{1}{2}\lambda_{SH}(v_H^2 + 2v_H h + hh)SS^\dagger \cdot$$

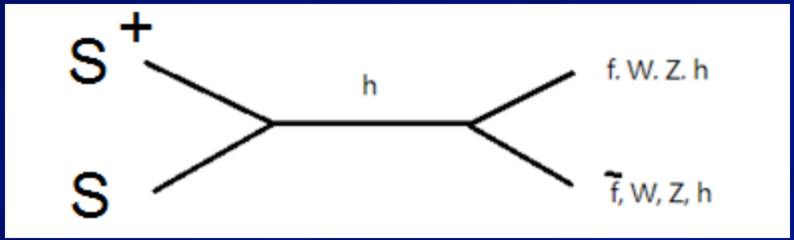
The first term will modify the mass of S from μ_S^2 to $M_D^2 = \mu_S^2 + \lambda_{SH} v_H^2/2$.

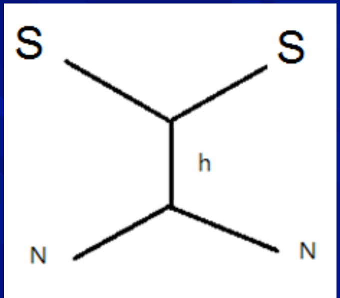


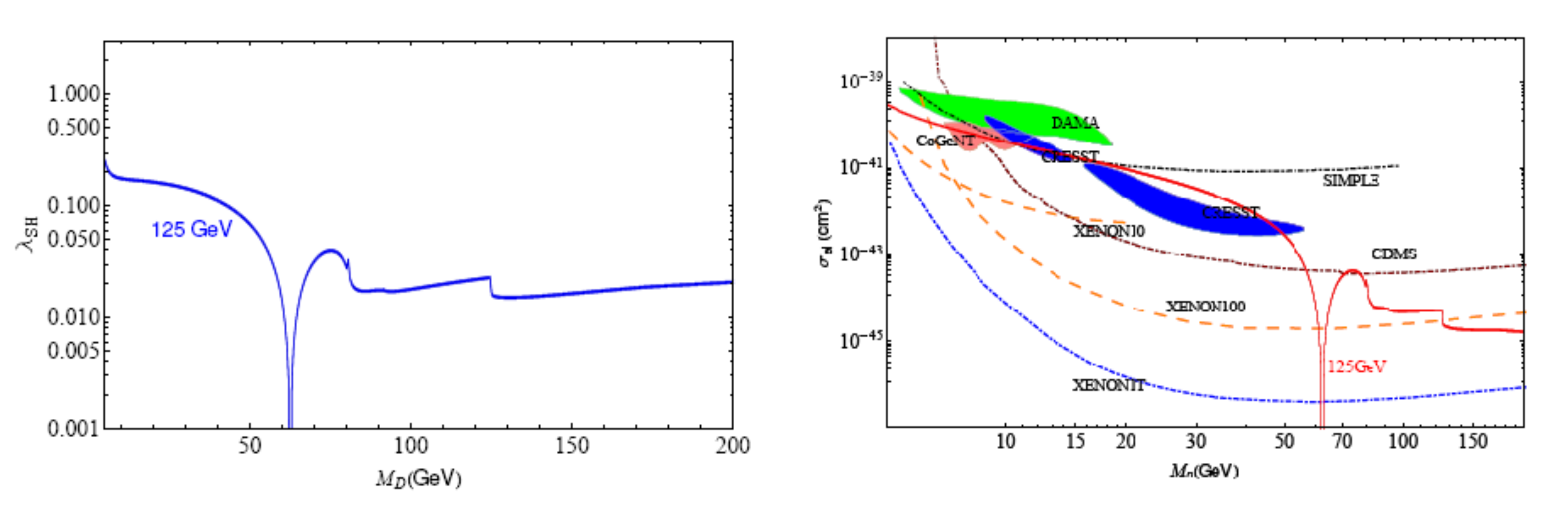


Real and imaginary parts of S have degenerate mass.

Higgs doublet tree level couplings to SM particles not modified. h is almost SM Higgs. Take m_h = 125 GeV for analysis.





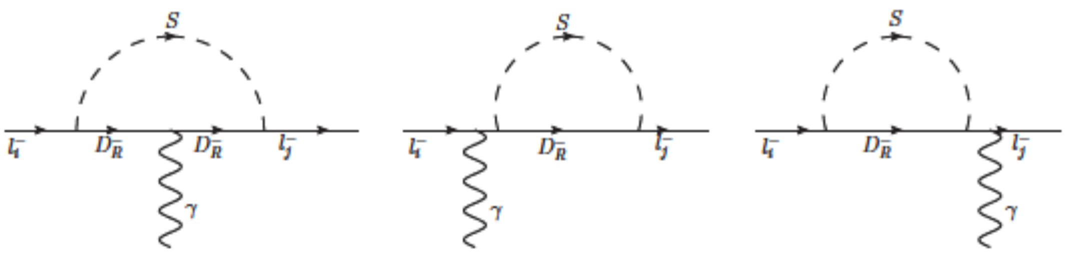


Consistent with all data, (left: relic density $\Omega h^2 = 0.11$, right: direct search) if dark matter is about half of Higgs mass or larger than 130 GeV

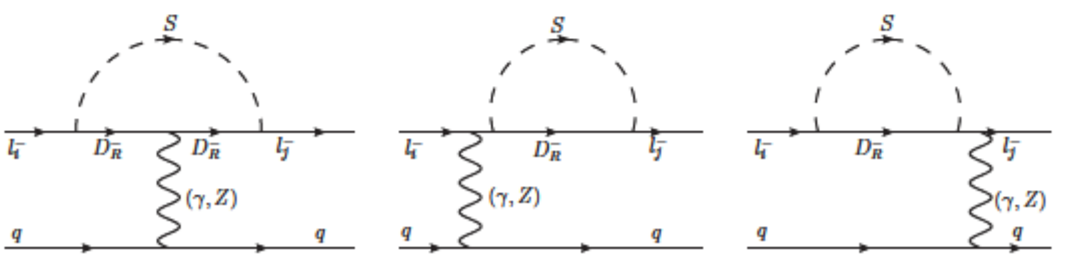
4. Leptonic FCNC and h ->γγ

At one loop level exchange D and S in the loop

μ -> e γ and μ – e conversion will be induced



Feynman diagram for $l_i \to l_j \gamma$.



Feynman diagrams for $l_i - l_j$ conversion.

$$\mathcal{L} = - \bar{l}_j \sigma^{\mu\nu}(A_{Lji}P_L + A_{Rji}P_R)l_i F_{\mu\nu} + [\sum_q eQ_q \bar{q}\gamma^\mu q \bar{l}_j B_{Lji}\gamma_\mu P_L l_i + H.c.] \, ,$$

and the functions $A_{L,R}$ and B_L are given by

$$A_{Lji} = Y_{Djk}Y^*_{Dki}\frac{e}{32\pi^2}\frac{1}{m_S^2}F_D(\frac{M_k^2}{m_S^2})m_i \, , \quad A_{Rji} = \frac{m_j}{m_i}A_{Lji},$$

$$B_{Lji} = Y_{Djk}Y^*_{Dki}\frac{e}{16\pi^2}\frac{1}{m_S^2}G_D(\frac{M_k^2}{m_S^2}) \, ,$$

$$F_D(z) = \frac{z^2 - 5z + 2}{12(z-1)^3} + \frac{z\ln z}{2(z-1)^4} \, ,$$

$$G_D(z) = \frac{7z^3 - 36z^2 + 45z - 16 + 6(3z-2)\ln z}{36(1-z)^4} \, .$$

The LFV $\mu \to e\gamma$ decay branching ratio is easily evaluated by

$$B(\mu \to e\gamma) = \frac{48\pi^2}{G_F^2 m_\mu^2}(|A_L|^2 + |A_R|^2).$$

The strength of $\mu - e$ conversion is measured by the quantity,

$$B^A_{\mu\to e} = \Gamma^A_{conv}/\Gamma^A_{capt} = \Gamma(\mu^- + A(N,Z) \to e^- + A(N,Z))/\Gamma(\mu^- + A(N,Z) \to \nu_\mu + A(N+1, Z-1).$$

$$\frac{B^A_{\mu\to e}}{B(\mu\to e\gamma)} = R^0_{\mu\to e}(A)\left|1 + \frac{\tilde{g}^{(p)}_{LV}V^{(p)}(A)}{A_R D(A)} + \frac{\tilde{g}^{(n)}_{LV}V^{(n)}(A)}{A_R D(A)}\right|^2 \, , \quad R^0_{\mu\to e}(A) = \frac{G_F^2 m_\mu^5}{192\pi^2\Gamma^A_{capt}}|D(A)|^2$$

and $\tilde{g}^{(p)}_{LV} = 2g_{LV(u)} + g_{LV(d)}, \tilde{g}^{(n)}_{LV} = g_{LV(u)} + 2g_{LV(d)}, g_{LV(q)} = -4eQ_q m_e B_L/(A_R)$.

The parameters $D(A)$, $V^{(p,n)}(A)$ are nuclei dependent quantities.

The current upper limit of branching ratio for this process is 2.4×10^{-12} at the 90% c.l.

We take $B(\mu \to e\gamma) = 1 \times 10^{-13}$ as the near future Experimental sensitivity will be improved.

There are several measurements of $\mu - e$ conversion on various nuclei.

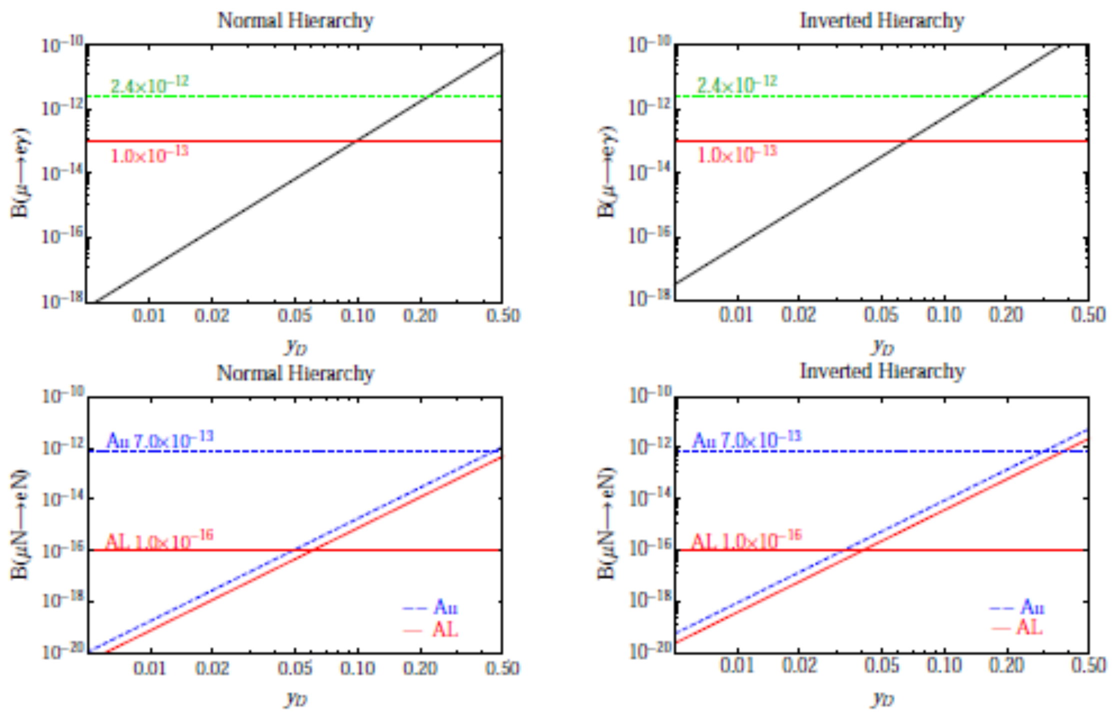
The best bound is for Au nuclei with the 90% c.l. experimental bound given by

$B^{\mathrm{Au}}_{\mu \to e} < 7 \times 10^{-13}$

$D(\mathrm{Au}) = 0.189,\ V^{(p)}(\mathrm{Au}) = 0.0974, V^{(n)}(\mathrm{Au}) = 0.146$ and $R^{0}_{\mu\to e}(\mathrm{Au}) = 0.0036$[25].

The sensitivities are expected to reach 10^{-16} for $\mu - e$ conversion using Al.

$D(\mathrm{Al}) = 0.0362,\ V^{(p)}(\mathrm{Al}) = 0.0161, V^{(n)}(\mathrm{Al}) = 0.0173$ and $R^{0}_{\mu\to e}(\mathrm{Al}) = 0.0026$



$B(\mu \to e\gamma)$(up) and $\mu - e$ conversion rate(down) depended on y_D with current(dashed line) and future(solid line) experiment constraints for normal(left) and inverted(right) hierarchy

h -> γγ

There are additional contribution by exchange Δ in the loop

The couplings of h to $\Delta^{-,--}$ come from $\lambda^{1,2}_{\Delta H}(\Delta^\dagger \Delta H^\dagger H)_{1,2}$ after H develops vev. They are given by

$$L \sim [\lambda^1_{\Delta H}(\Delta^+\Delta^- + \Delta^{++}\Delta^{--}) + \lambda^2_{\Delta H}(\Delta^{++}\Delta^{--} + \frac{1}{2}\Delta^+\Delta^-)]vh\,.$$

Combining with contributions from W and top in the loop, the $h \to \gamma\gamma$ rate is modified by a factor $R_{\gamma\gamma} = \Gamma(h \to \gamma\gamma)_{U(1)_D}/\Gamma(h \to \gamma\gamma)_{SM}$ given by

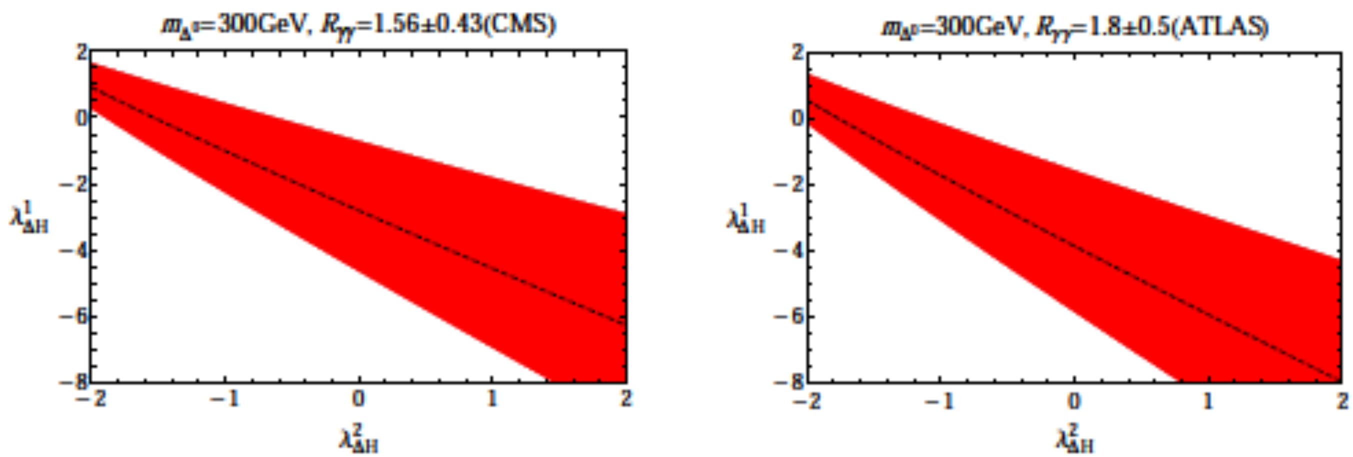
$$R_{\gamma\gamma} = |1 + \frac{v^2}{2}\frac{1}{A_1(\tau_W) + N_c Q_t^2 A_{1/2}(\tau_t)}\{\frac{\lambda^1_{H\Delta} + \frac{1}{2}\lambda^2_{H\Delta}}{m^2_{\Delta^-}}A_0(\tau_{\Delta^-}) + \frac{4(\lambda^1_{H\Delta} + \lambda^2_{H\Delta})}{m^2_{\Delta^{--}}}A_0(\tau_{\Delta^{--}})\}|^2$$

where $\tau_i \equiv (m_h^2/4m_i^2)$. $A_1(\tau_W)$ and $A_{1/2}(\tau_t)$ come from SM W boson and top quark contributions. $A_0(\tau_\Delta)$ comes from new scalars in the model. They are given by

$$A_0(x) = -x^{-2}[x - f(x)];\ A_{1/2}(x) = 2x^{-2}[x + (x-1)f(x)];$$
$$A_1(x) = -x^{-2}[2x^2 + 3x + 3(2x-1)f(x)];$$
$$f(x) = \begin{cases} \arcsin^2\sqrt{x}, & x \geq 1 \\ -\frac{1}{4}[\ln\frac{1+\sqrt{1-x^{-1}}}{1-\sqrt{1-x^{-1}}} - i\pi]^2, & x < 1 \end{cases}$$

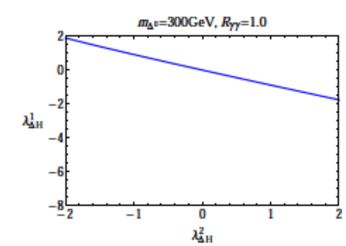
$$m^2_{\Delta^0} = \mu^2_\Delta + \frac{1}{2}\lambda^1_{H\Delta}v^2\,,$$
$$m^2_{\Delta^-} = \mu^2_\Delta + \frac{1}{2}\lambda^1_{H\Delta}v^2 + \frac{1}{4}\lambda^2_{H\Delta}v^2\,,$$
$$m^2_{\Delta^{--}} = \mu^2_\Delta + \frac{1}{2}\lambda^1_{H\Delta}v^2 + \frac{1}{2}\lambda^2_{H\Delta}v^2\,,$$

To enhance the ratio $R_{\gamma\gamma}$, negative $\lambda^{1,2}_{H\Delta}$ are preferred. With fixed m_{Δ^0}, negative $\lambda^2_{H\Delta}$ implies that $m_{\Delta^0} > m_{\Delta^-} > m_{\Delta^{--}}$. From Fig. 5, we can see that with negative $\lambda^{1,2}_{H\Delta}$ of order $O(1)$, the ATLAS and CMS results on $h \to \gamma\gamma$ can be reproduced. If one controls the magnitude of $\lambda^{1,2}_{H\Delta}$ as small as possible from perturbetivity consideration, the optimal values for $\lambda^{1,2}_{H\Delta}$ are around -0.8 and -0.6, respectively. With these values, Δ^- and Δ^{--} masses are given by 284.5 GeV and 268 GeV. With the same parameters, the predicted value for $R_{\gamma Z}$ are shown in Fig.

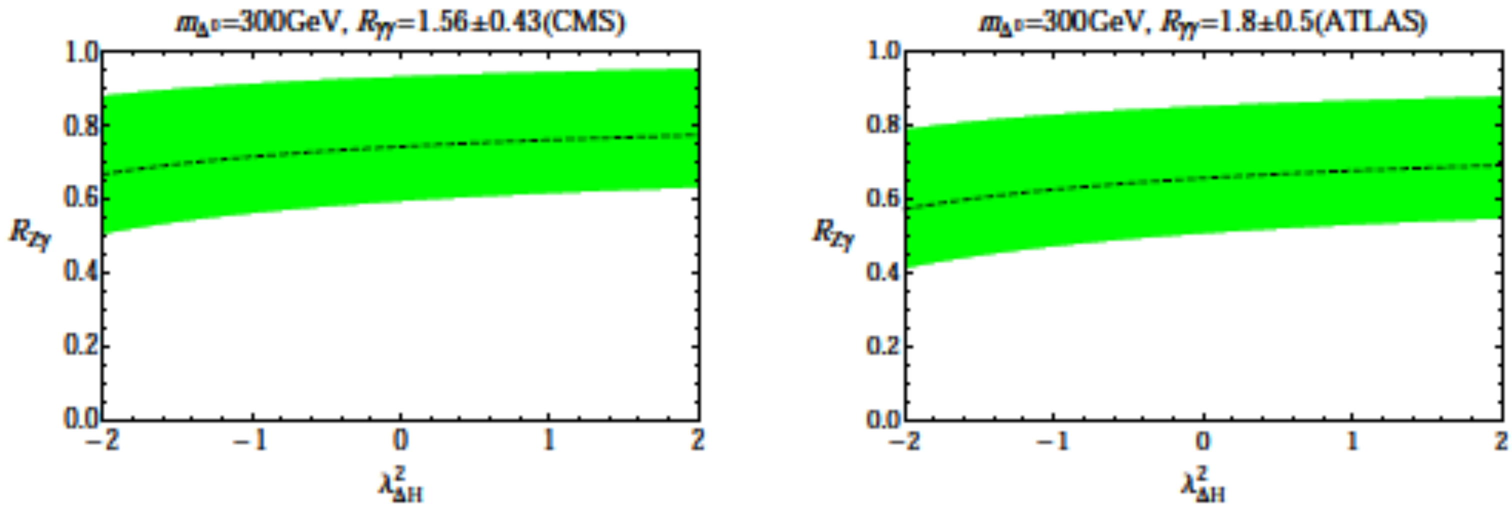


Constraints on $\lambda^1_{H\Delta}$ and $\lambda^2_{H\Delta}$ with $m_{\Delta^0} = 300\text{GeV}$.

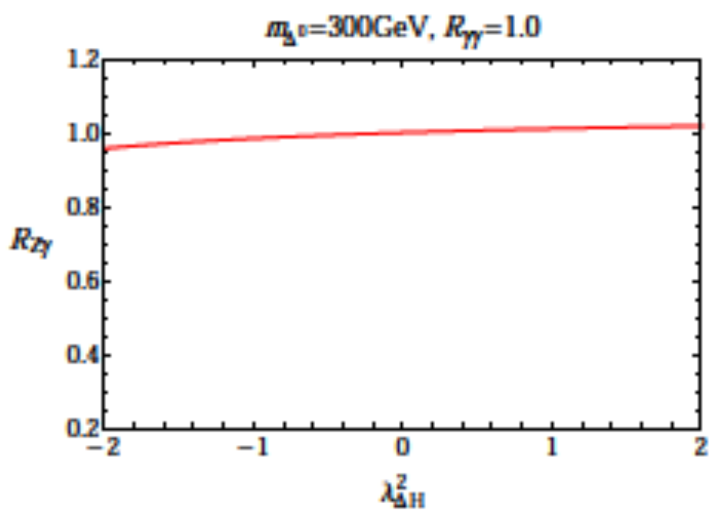
With $\lambda^2_{H\Delta} > 0$, the mass hierarchy for the component fields in the triplet is $m_{\Delta^0} < m_{\Delta^-} < m_{\Delta^{--}}$. In this case, the new contributions may cancel out if $\lambda^1_{H\Delta}$ is kept negative. We demonstrate this possibility in Fig. , where $R_{\gamma\gamma}$ is kept to be 1.



Constraints on $\lambda^1_{H\Delta}$ and $\lambda^2_{H\Delta}$ with $R_{\gamma\gamma} = 1$ for $m_{\Delta^0} = 300\text{GeV}$.



Shift factor of the $h \to Z\gamma$ with the same parameters for $h \to \gamma\gamma$.



Shift factor of the $h \to Z\gamma$ with the same parameters for $R_{\gamma\gamma} = 1$.

5. Conclusions

Possible to construct models with two loop radiative inverse seesaw neutrino masses and at the same time with viable dark matter candidate.

Due to large Yukawa couplings, there are large testable leptonic FCNC effects in $\mu \to e\,\gamma$ and μ –e conversion

Exchange of Δ particle in the loop, enhanced $h \to \gamma\gamma$ reported by LHC can be produced.