ASTROPARTICLES

ESIPAP - 2014

François Montanet

Plan of the course

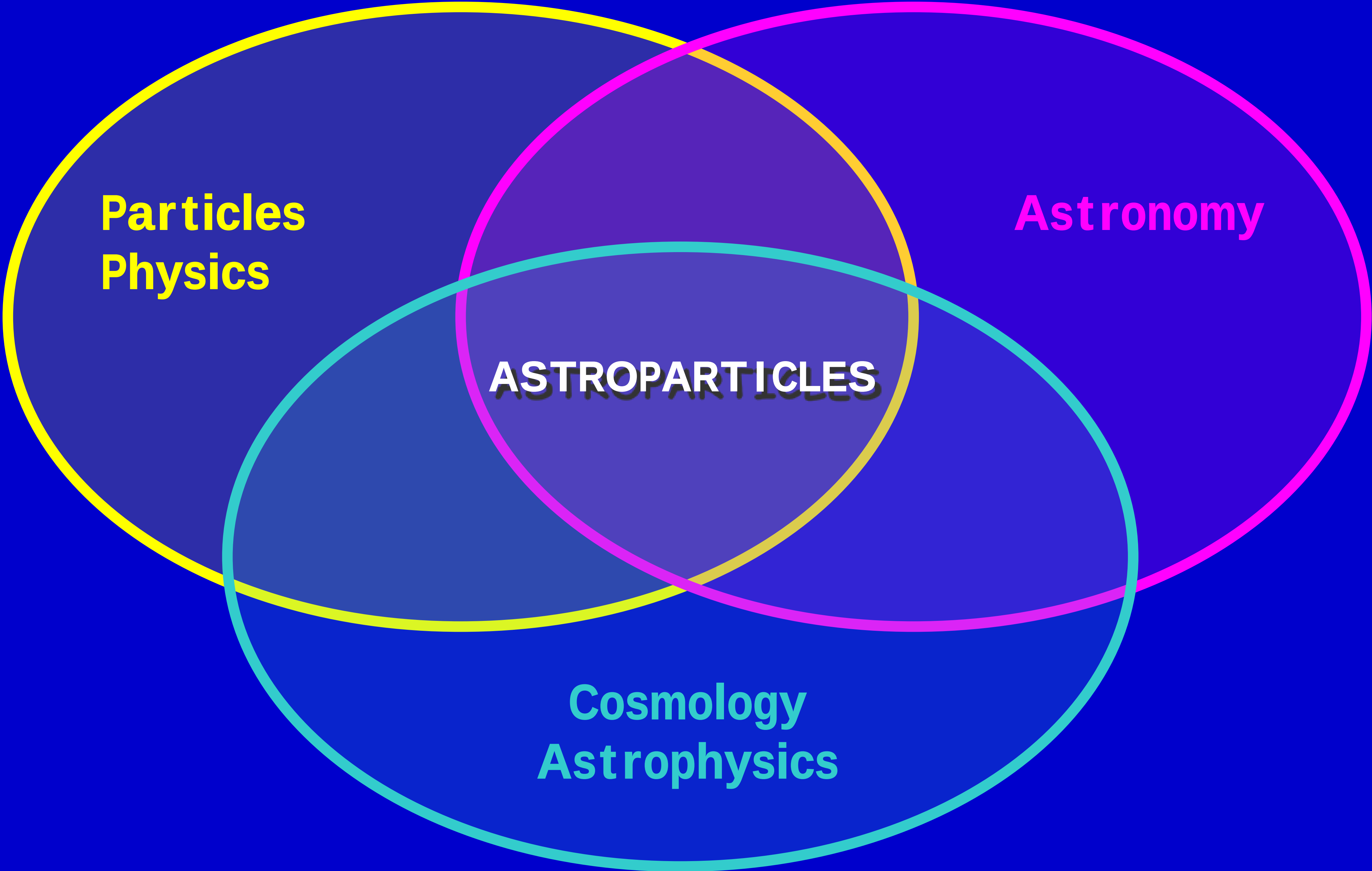
- Introduction & a very brief History
- Nature and properties of Astroparticles
- Propagation medium, IGM, ISM and atmosphere
- Astrophysical Sources
  - Astrophysical shocks
  - Fermi acceleration
  - Standard Model for the production of galactic CR, SNR
  - Gamma-ray sources, pulsars
  - AGN and other extragalactic sources
  - Neutrinos sources
- Cosmological Sources
  - Cosmological implications
  - Dark Matter as a source of CR
  - "top-down" type of sources at UHE
- Propagation
  - CR propagation in the Galaxy: The Leaky box model
  - VHE g-rays propagation
  - UHECR propagation
- Observables & Observations
  - Primary CR detection (on top of atmosphere)
  - Air Showers Development Models
  - Gamma-ray (EM) induced showers detection
  - Hadronic Showers Models and Detection
  - UHECR detection
- Dark matter search
- Neutrino Physics with astroparticules

Why studying Astroparticles

Open questions

Understand our Universe at extreme scales

- The Higgs boson or the origin of mass
- Nature and mass of Neutrinos
- Fundamental symmetries: CP , supersymmetry
- New dimensions in physical space ?
- What is our Universe made of ?
- Sources and propagation of cosmic rays ?
- New Physics at E >> E(LHC) ?


Particles
Physics
Astronomy
ASTROPARTICLES
Cosmology
Astrophysics

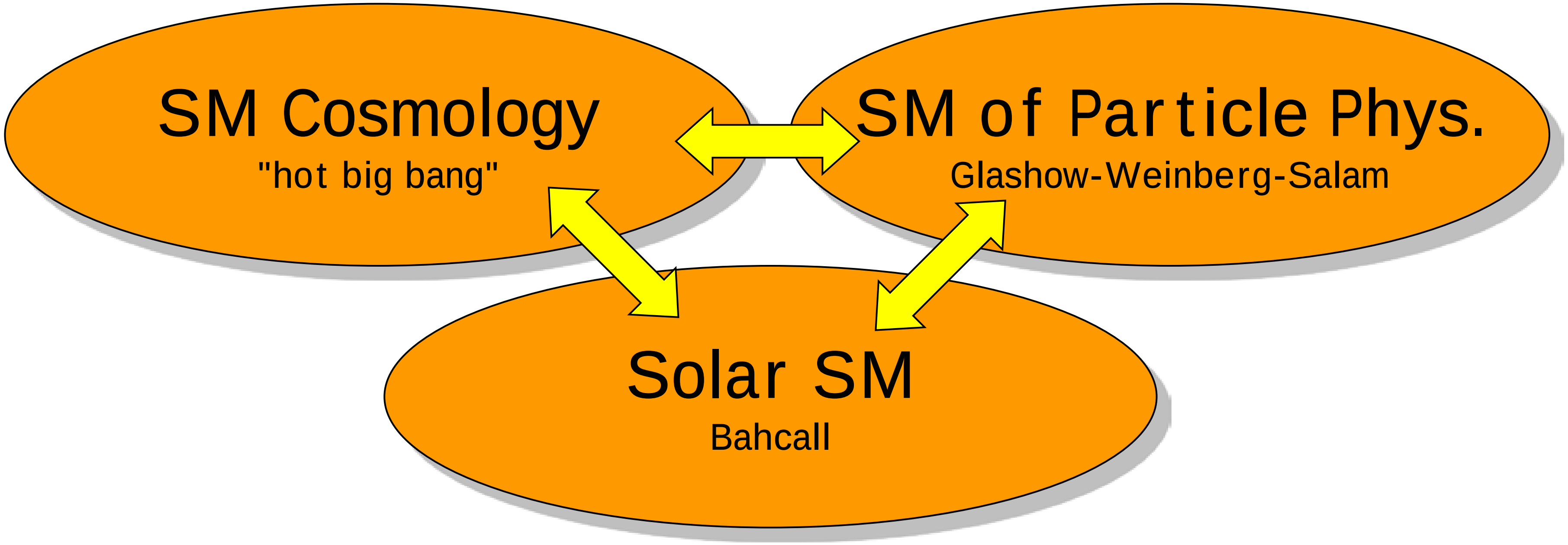
Astroparticles & HEP

Accelerator Physics

Cosmology and Astroparticle physics

Search for new physics, beyond the Standard Model of particle physics.

Matching "standard models"



Examples of happy unions...

Nucleosynthesis $\Rightarrow N_\nu$

$\Omega h^2, LSS... \Rightarrow \sum m_\nu < ...$

... as well as some disputes...

BigBang ↔ GWS ↔ Solar SM

Dark Matter
Dark Energy
Inflation
Matter Anti-Matter Asymmetry

Solar n deficit
Atmospheric n anomalies

Direct searches

↓

new particles production-observations

(Tevatron, LHC)

New Physics

This decade's grail: solve the puzzle of electroweak symmetry breaking

Indirect searches

$FCNC$, CP

$FV \rightarrow \mu \rightarrow e\gamma$

d_n^e

B physics

Progress in Theory

Supergravity

$\rightarrow$ Superstrings, M-Theory

Cosmology

Measure the parameters of the Univers and their evolution

Astroparticle Physics

Neutrino Physics

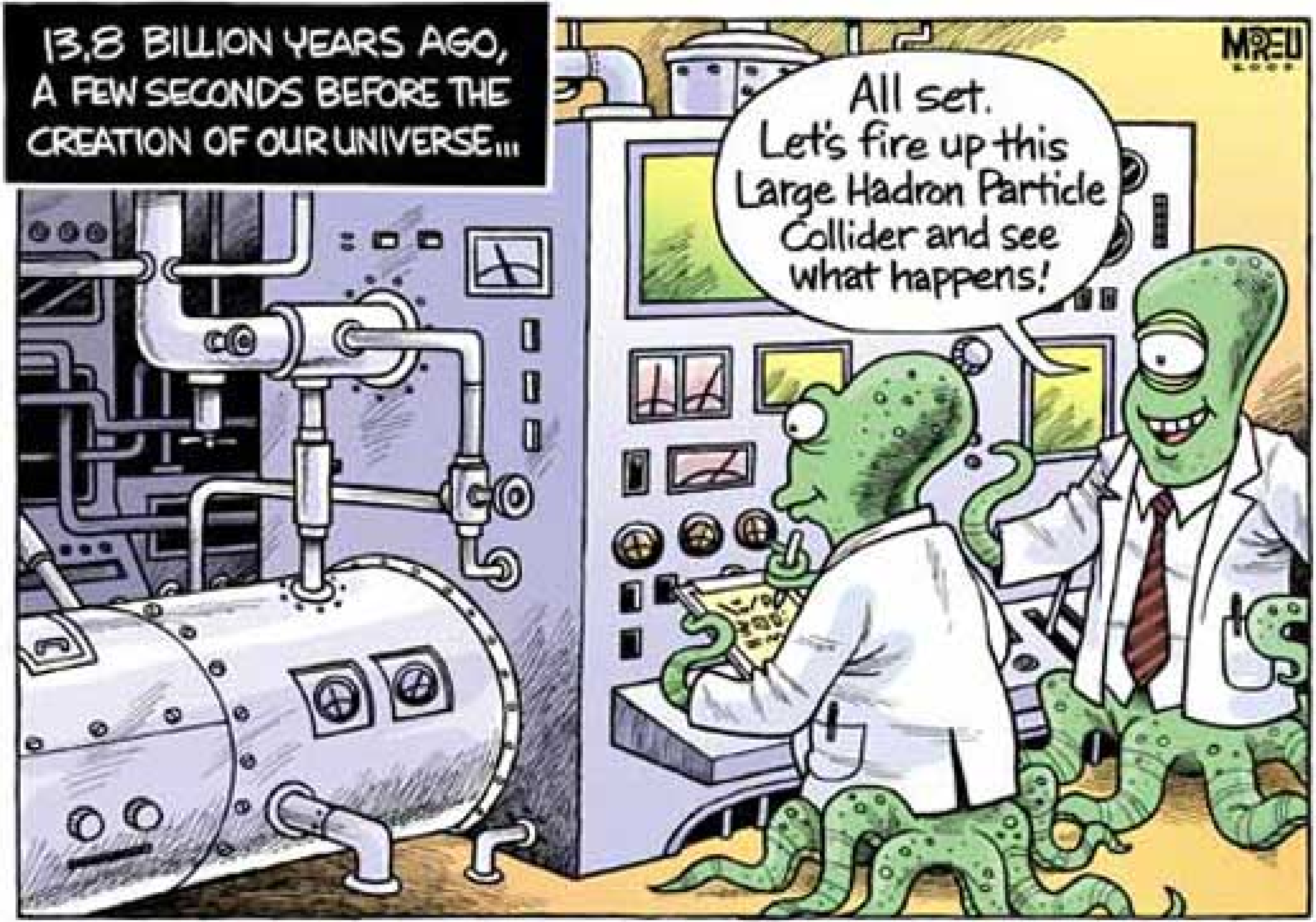
Cosmic Rays

γ Astronomy

Gravitationnal Waves

↓

New Physics probes


13,8 BILLION YEARS AGO, A FEW SECONDS BEFORE THE CREATION OF OUR UNIVERSE...
All set. Let's fire up this Large Hadron Particle Collider and see what happens!

What are "Astroparticles"

What we know

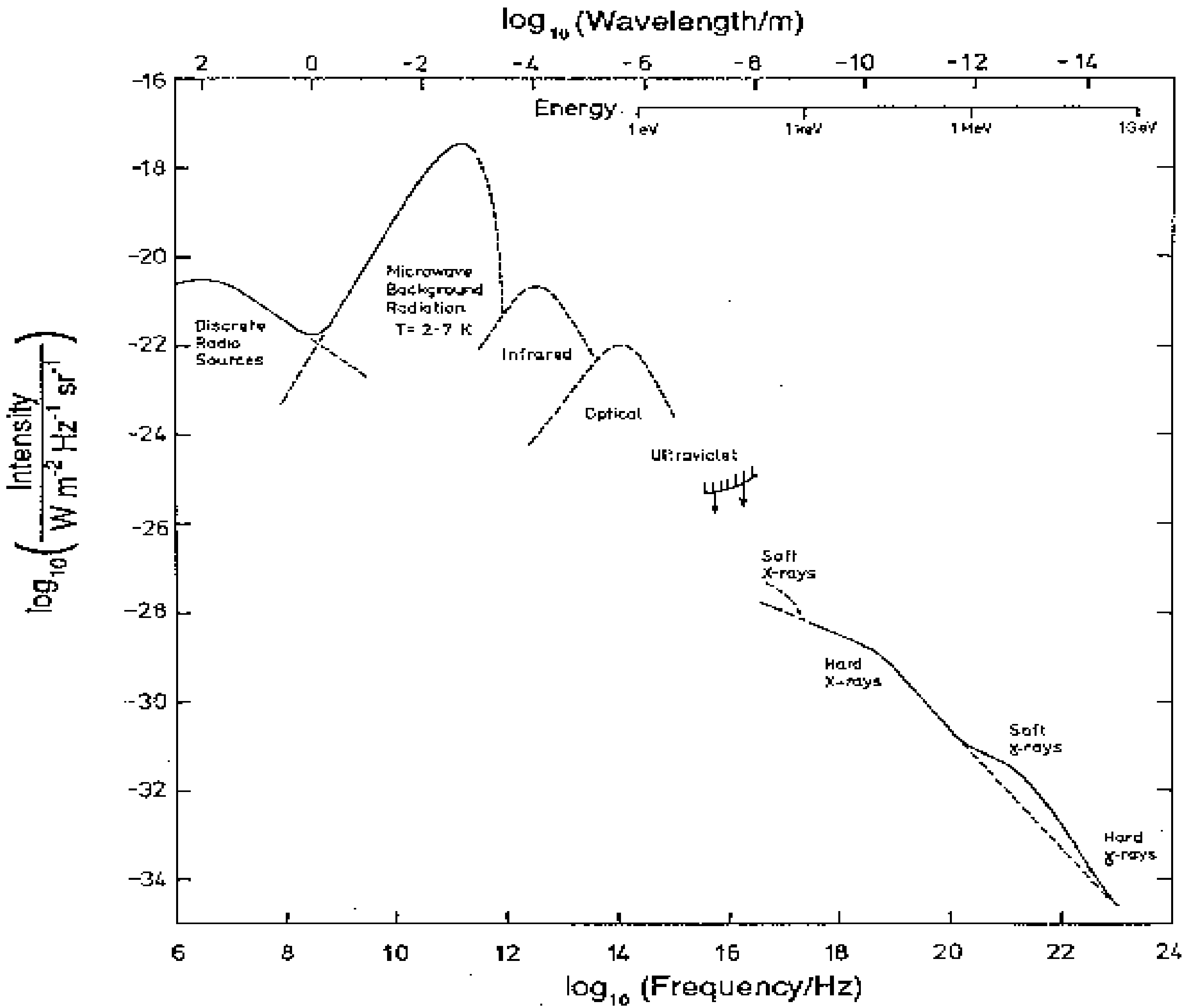
Let there be light !

All what we know in astrophycics is thanks to the light !

- Temperatures, stars masses, galaxies, magnetic fields, chemical composition, age of stars and structures…
- Nuclear reactions, galactic and extragalactic hydrodynamics, MHD, explosions, nucleosynthesis, past, future… EVERYTHING !

Well, almost everything…

A multi-wavelength sky

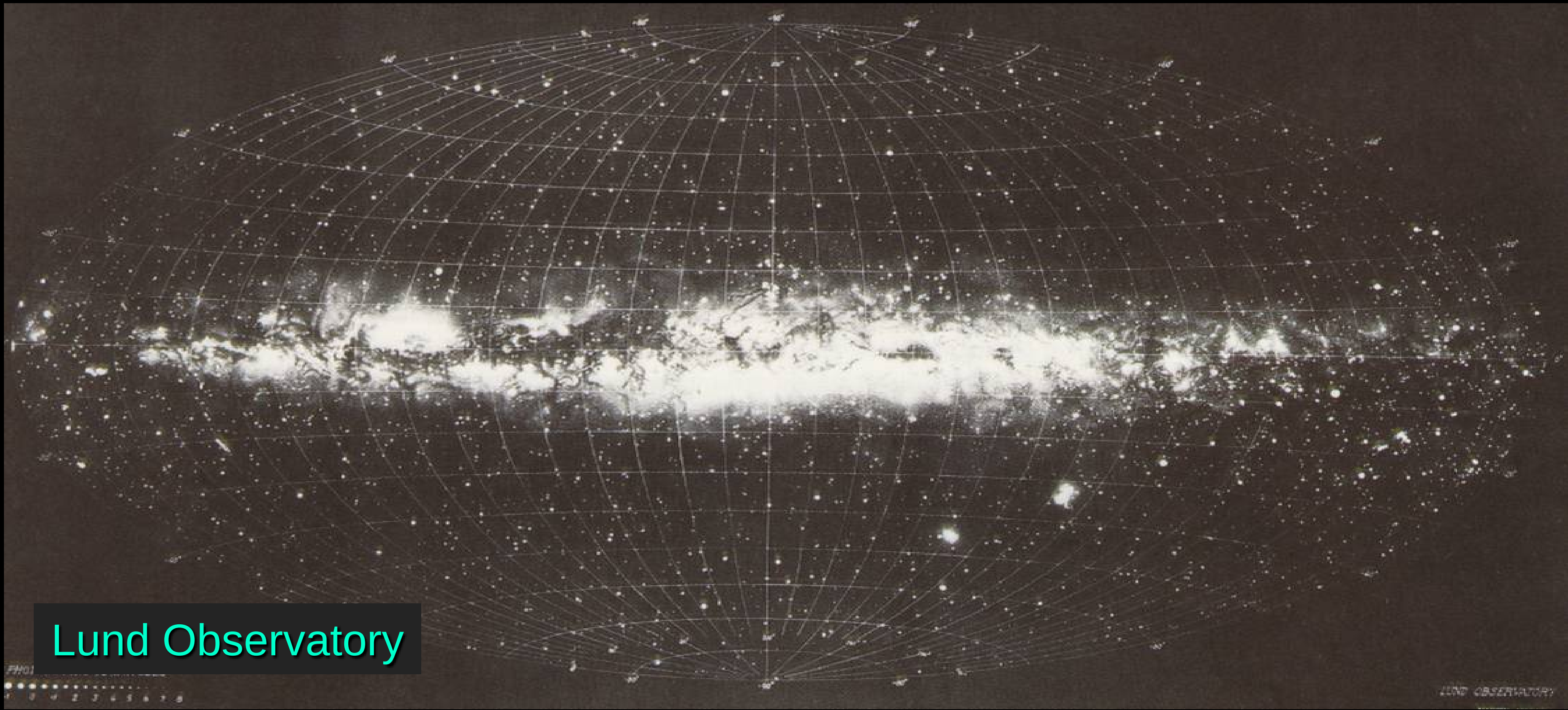


Our Galaxy

The optical Milky Way

Our Galaxy

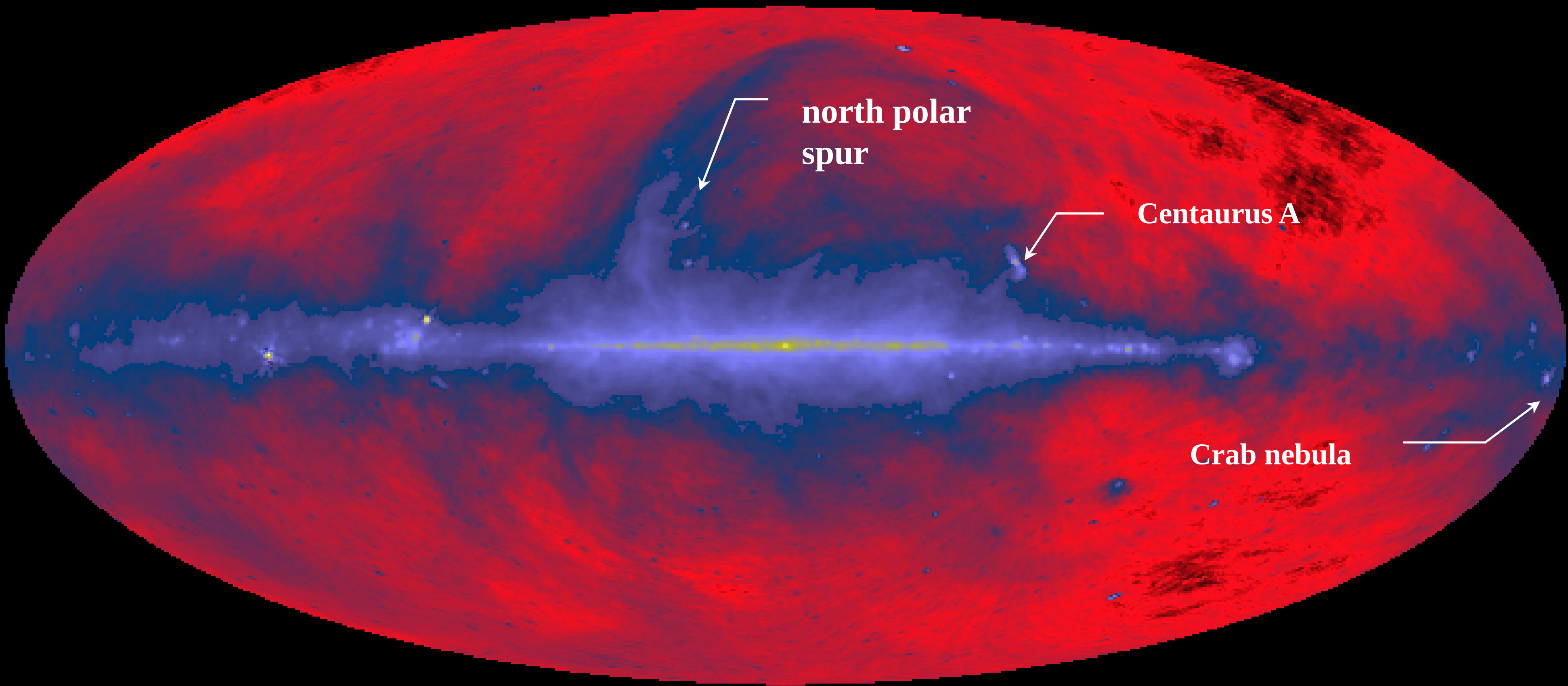
The optical Milky Way



Our Galaxy

The Milky Way : Radio at 73cm

408 MHz / 73.5 cm / 1.6 10^{-6} eV)

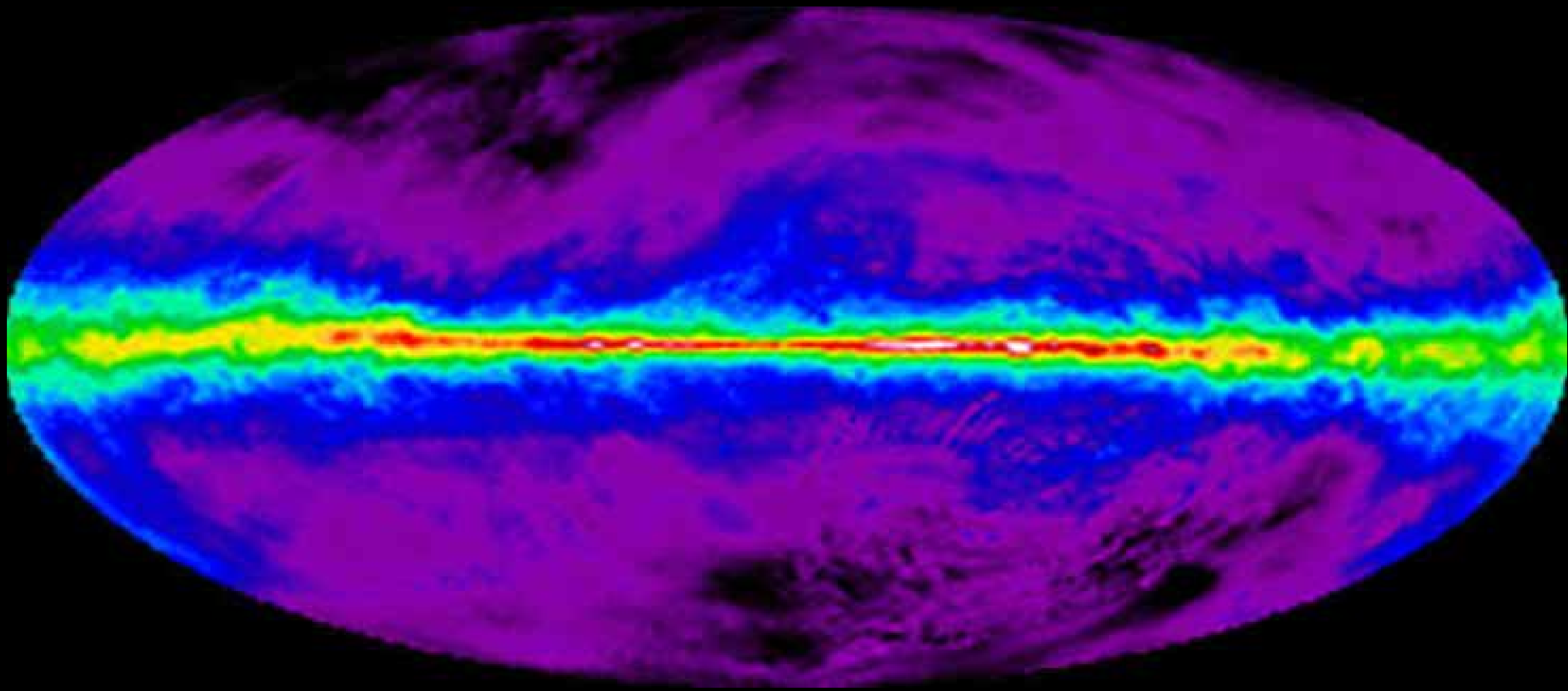


Essentially from the movement of ultra relativistic electrons probably issue from supernovae remnants in the galactic magnetic field.

Our Galaxy

The Milky Way : Radio at 21 cm

(~1.42 GHz / 21.1 cm / 5.9 10^{-6} eV)

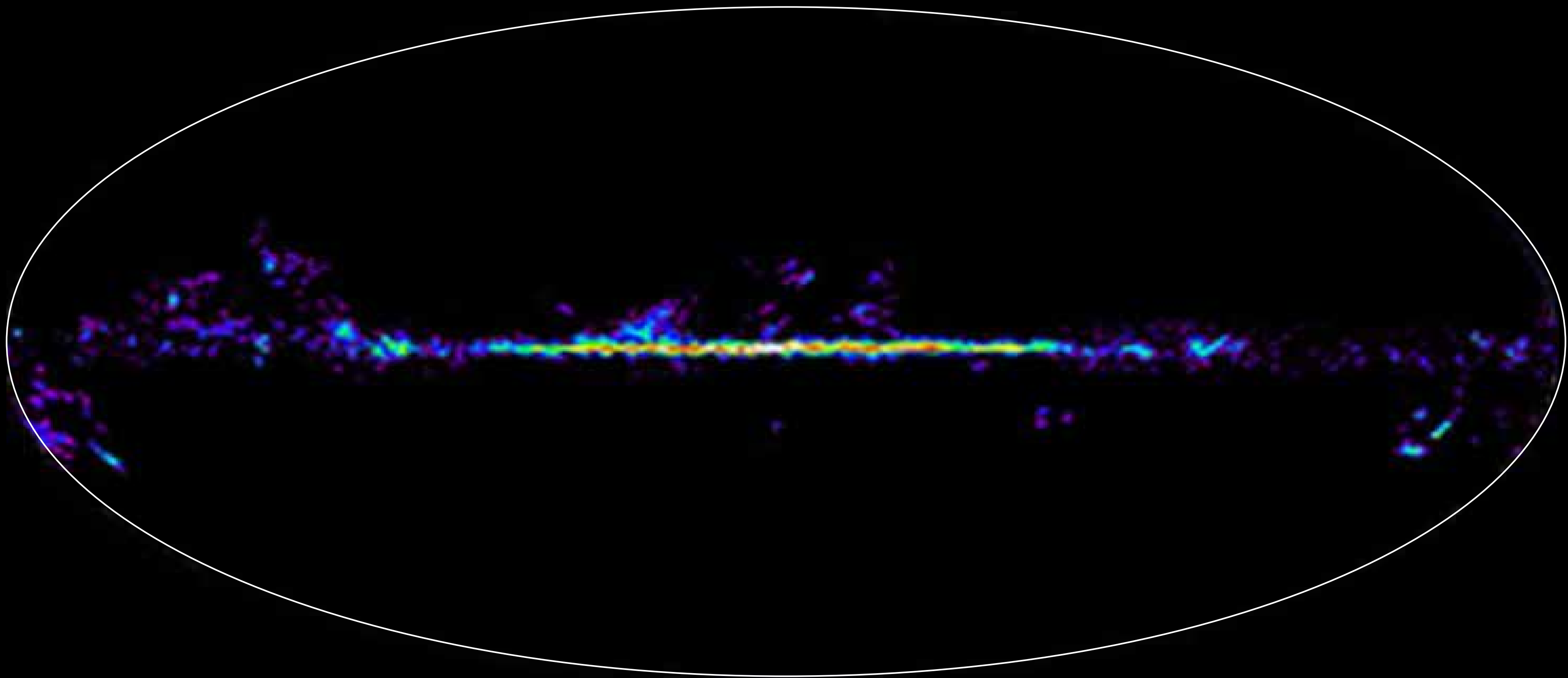


Hyperfine transition of hydrogen. Structures are due to the column density of atomic hydrogen clouds along the line of sight.

Our Galaxy

The Milky Way : Radio at 2,6mm

Millimetric waves (115 GHz / 2.6 mm / 4.7 10^{-4} eV)

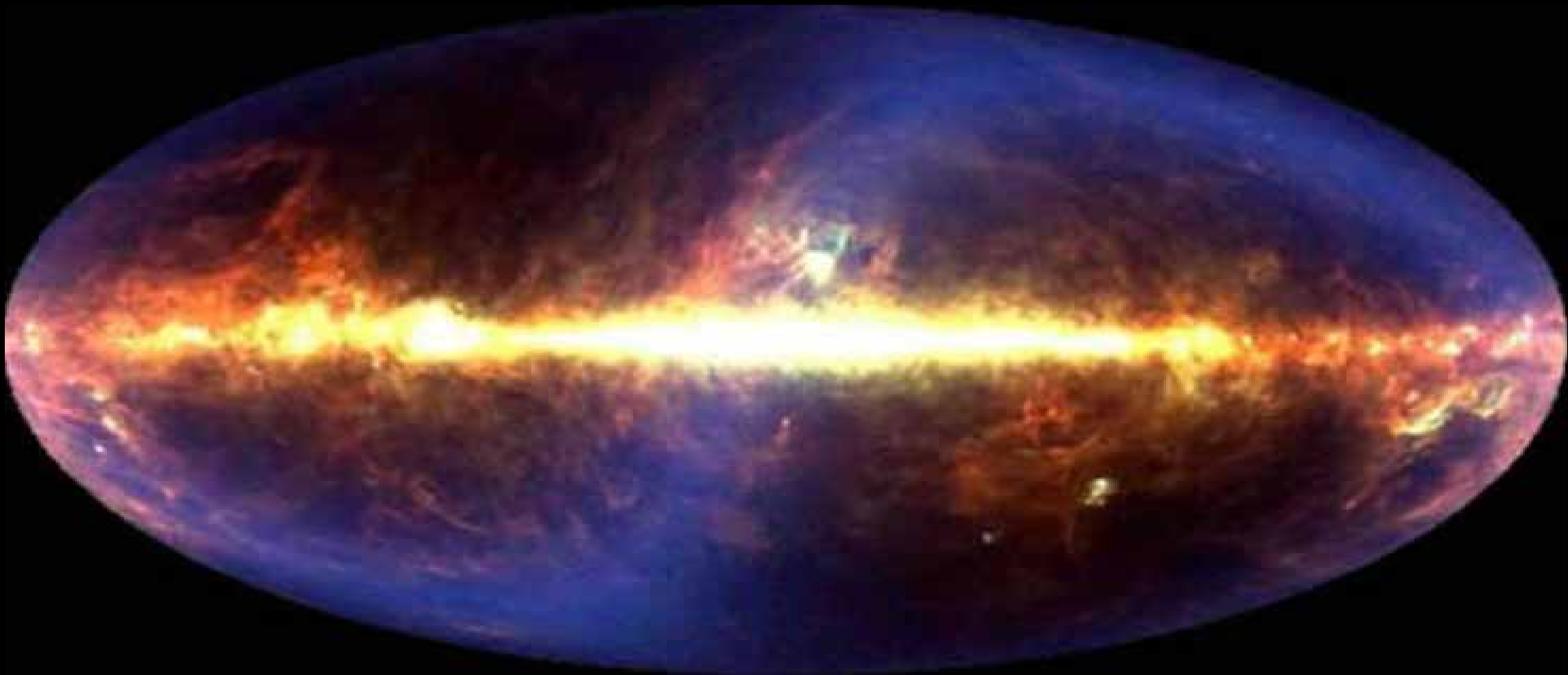


Rotation mode ray of carbon monoxide. One assumes that CO abundance is proportional to that of cold molecular hydrogen (directly undetectable).

Our Galaxy

Infra red

Infrared (3 10^3 to 25 10^3 GHz / 100 to 12 μm / 0.01 to 0.1 eV)

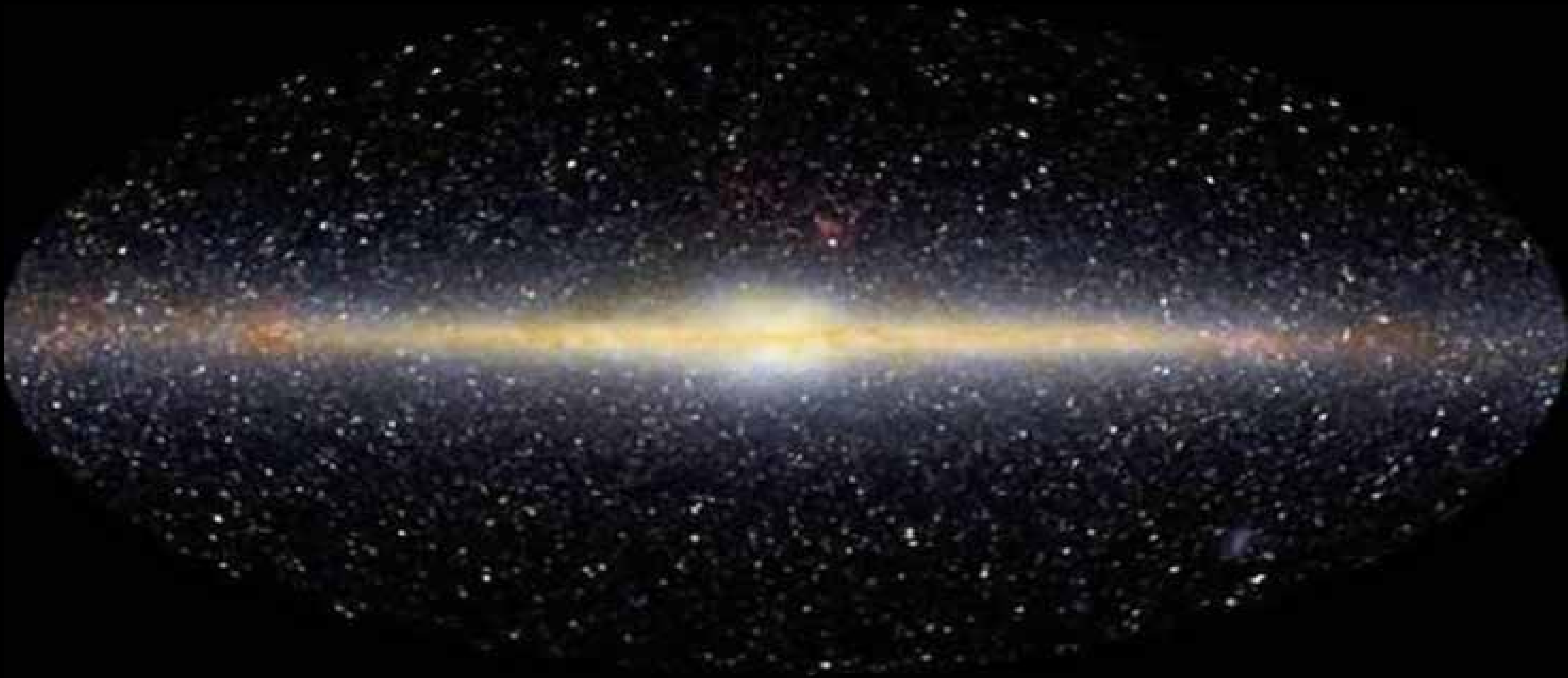


Thermal emission, due to interstellar dust heated by starlight.

Our Galaxy

It structure is clearly visible in IR (COBE satellite).

Near Infrared (86 10^3 à 240 10^3 GHz / 1.25 à 3.5 µm / 0.35 à 1 eV).



Giant stars emission in the disk and in the bulb

Our Galaxy

Optical

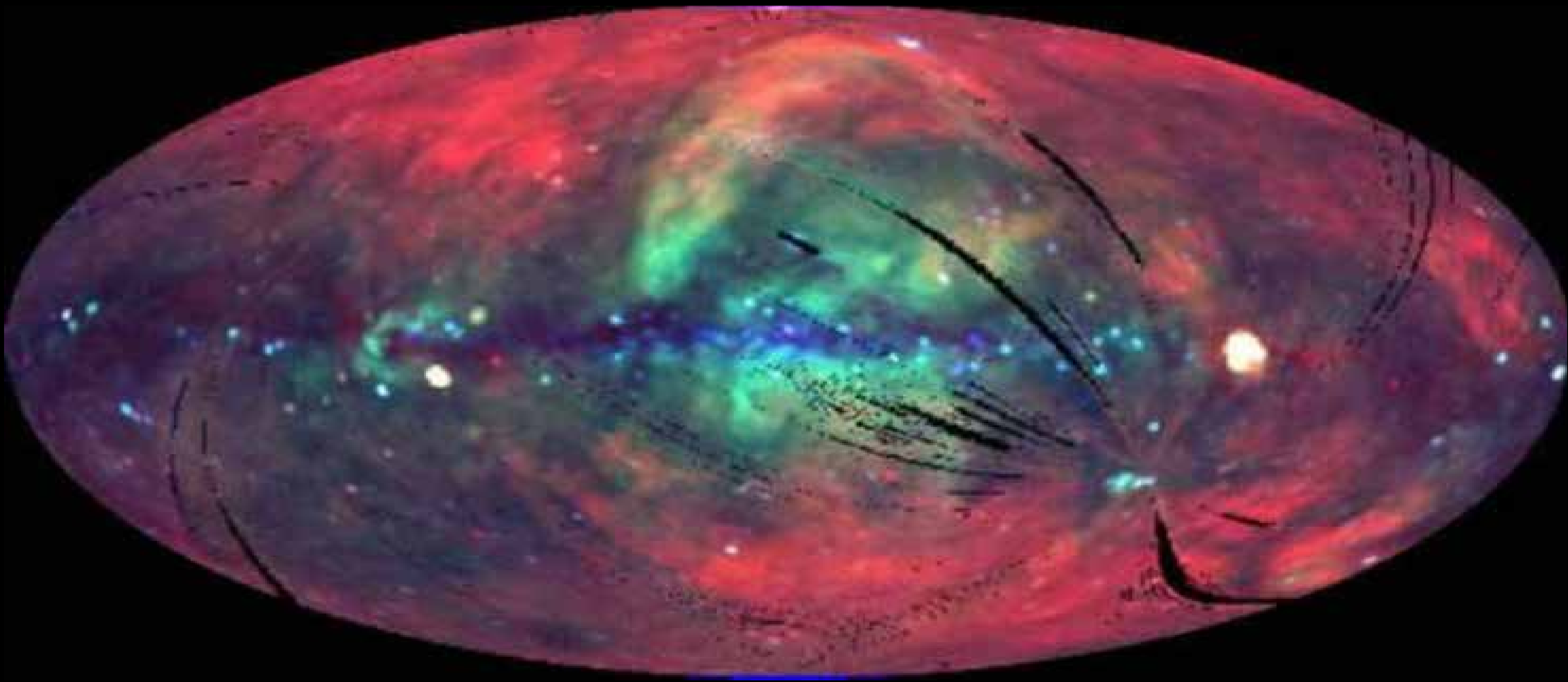
Visible (460 10^3 GHz / 0.65 μm / 2 eV – red)

Visible light is absorbed by interstellar dust clouds.
Only stars close enough to the solar system (few parsec) are seen.

Our Galaxy

X-rays

X -rays (60.10^6 to 360.10^6 GHz / 5 to 8.3 nm / 0.25 to 1.5 keV).

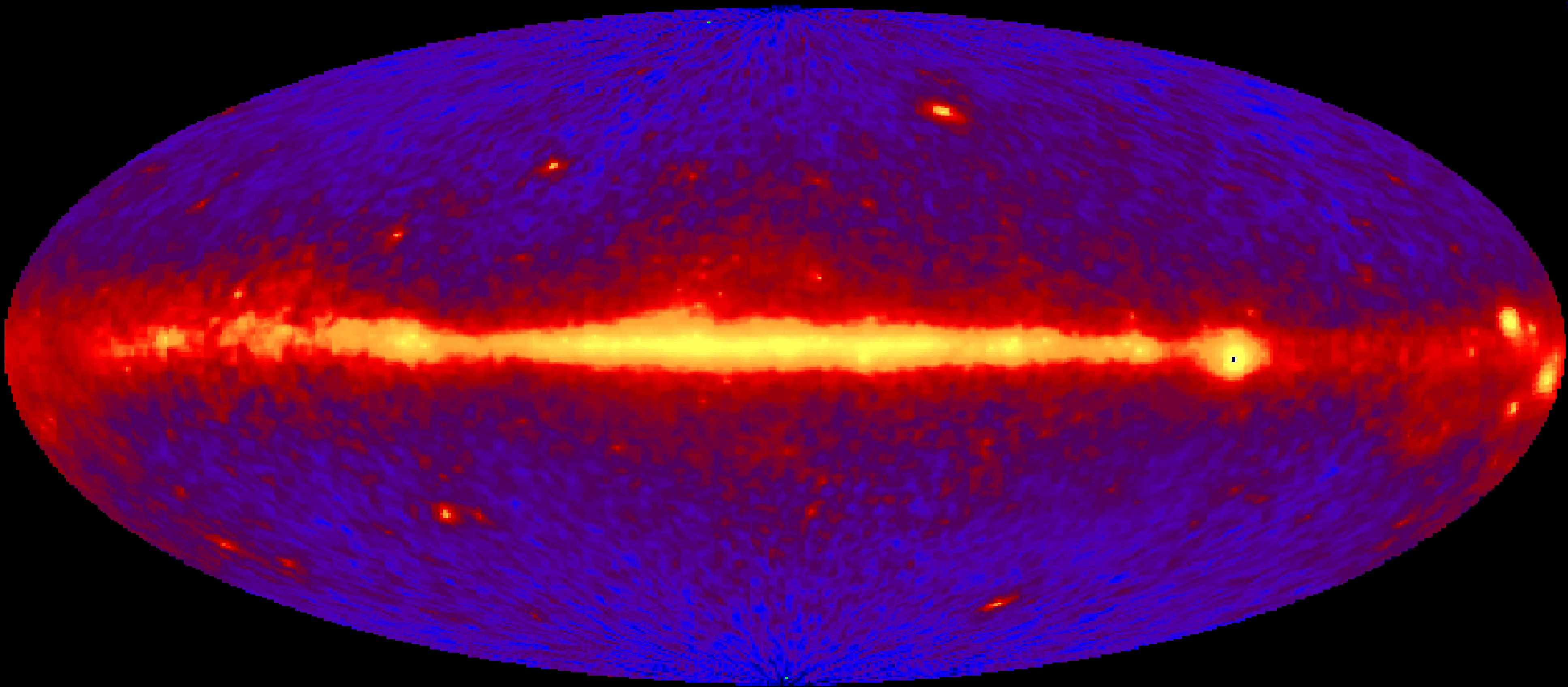


Diffuse X-ray emission from overheated and shocked gas.

Our Galaxy

Gamma-rays

Gamma-rays (> 2.4 10^{13} GHz / < 12.5 fm / >100 MeV).



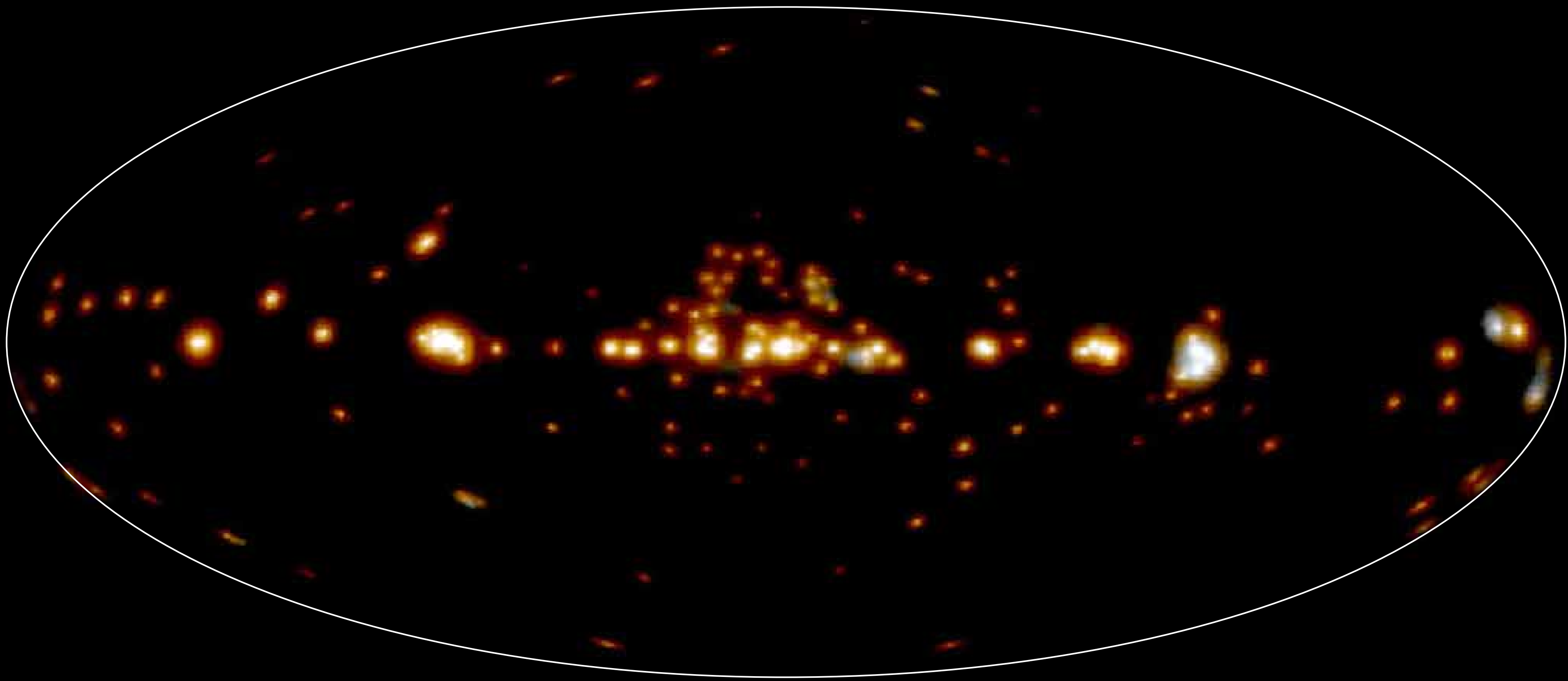
Photons (gammas) from the decay of neutral pions produced in the interaction of CR with interstellar matter, from the Bremsstrahlung of CR and from the inverse Compton of relativistic electrons with ambient photons.

Our Galaxy

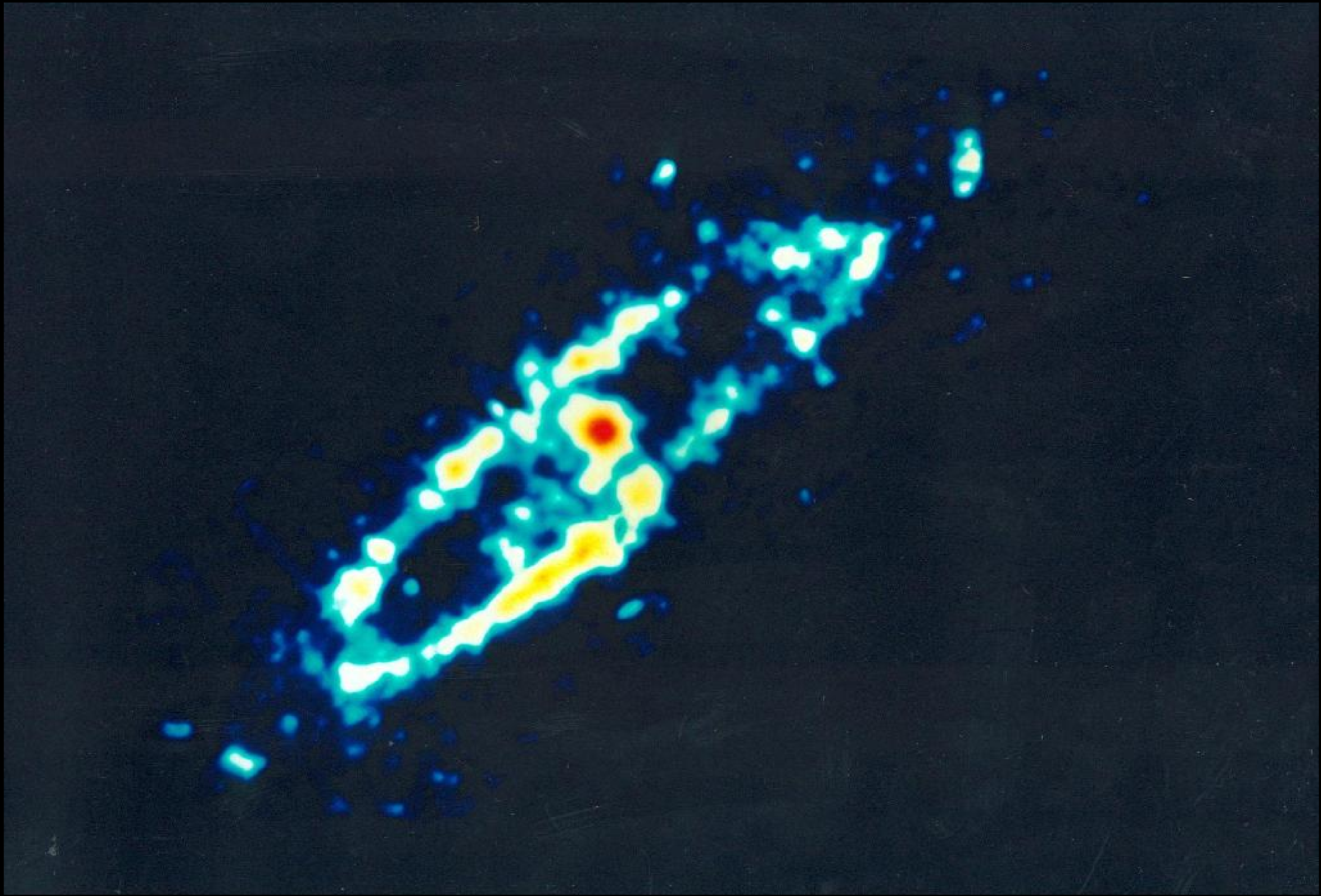
HE gamma-rays (>100MeV EGRET satellite)

Resolved point-like sources:

Binary systems, pulsars, SN remnants...



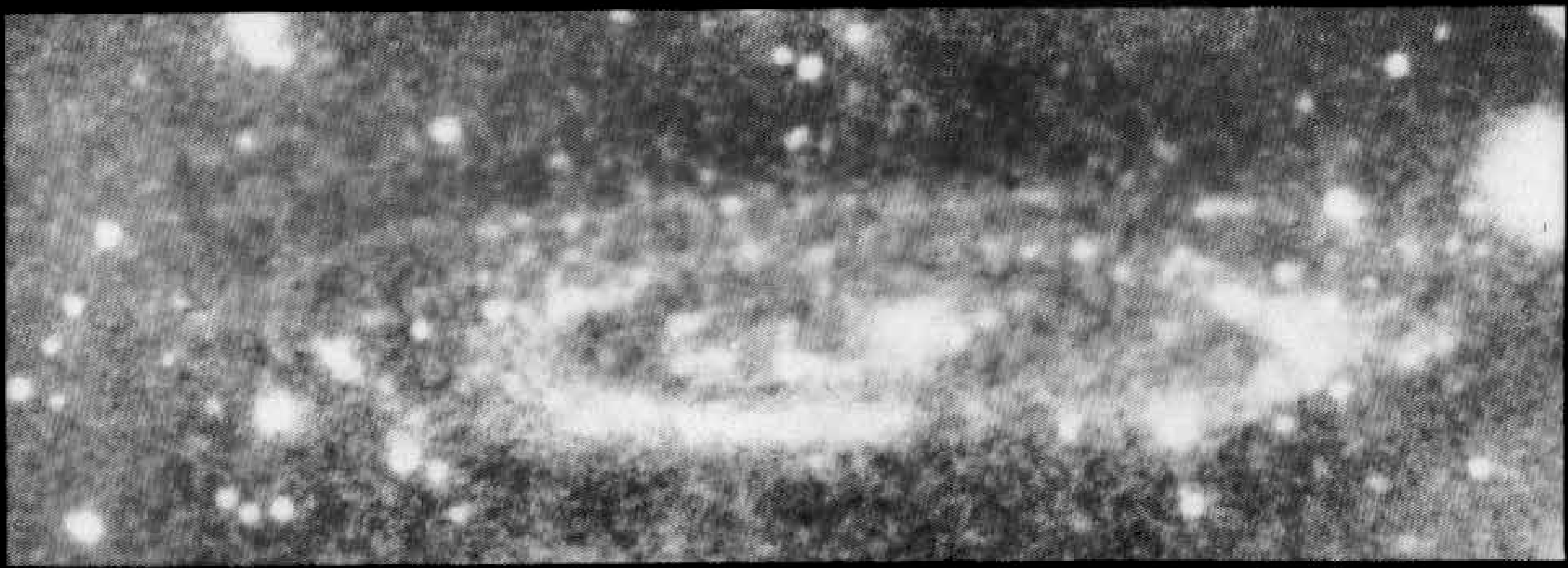
Andromeda (M31): IR



Star forming regions in spiral arms

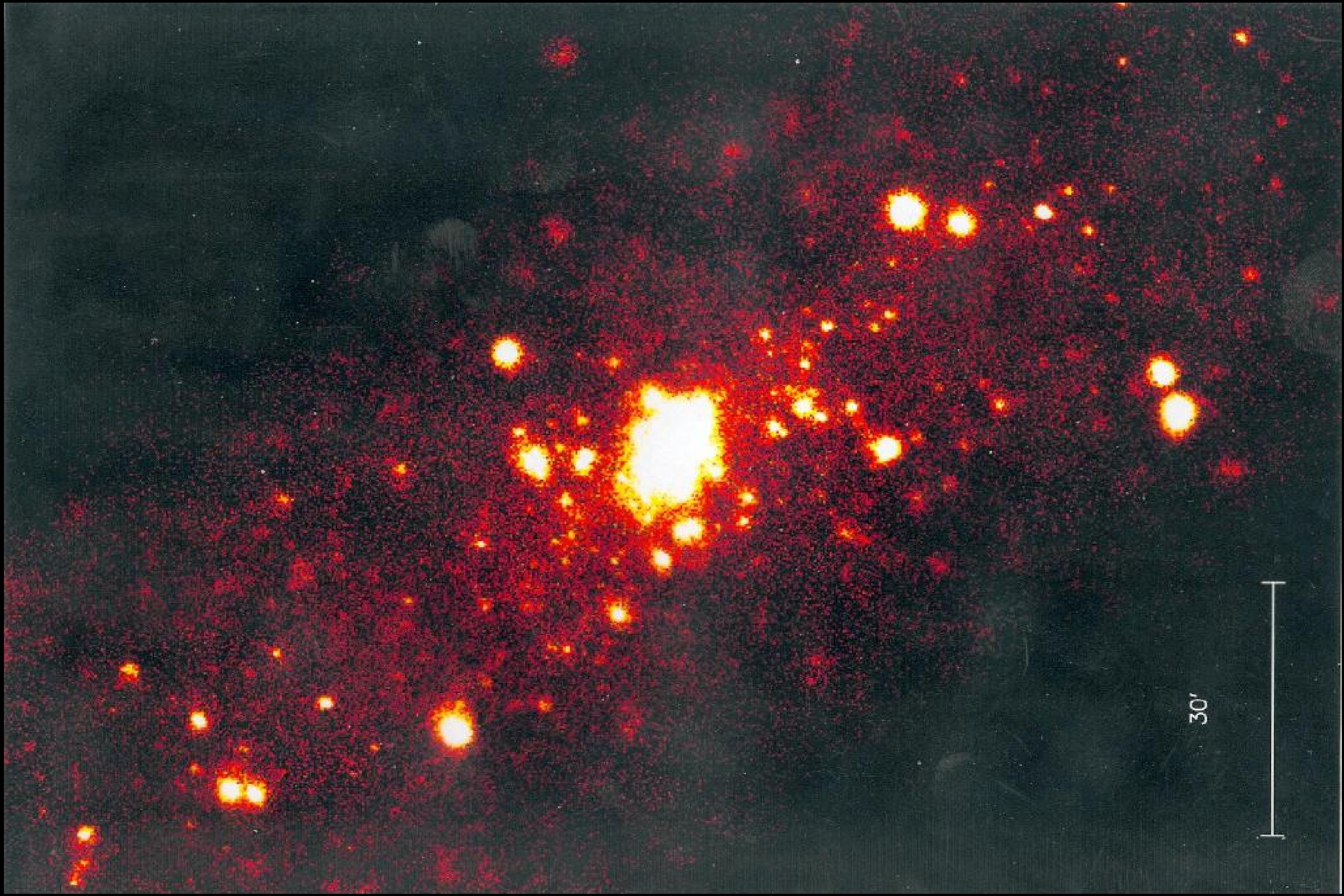
Andromeda (M31): UV

This photograph of the galaxy M31 reveals the prominence at ultraviolet wavelengths (2000 Å) of young stars in the spiral arms over the older population in the central bulge.
(B. Milliard/ Laboratoire d'Astronomie Spatiale).



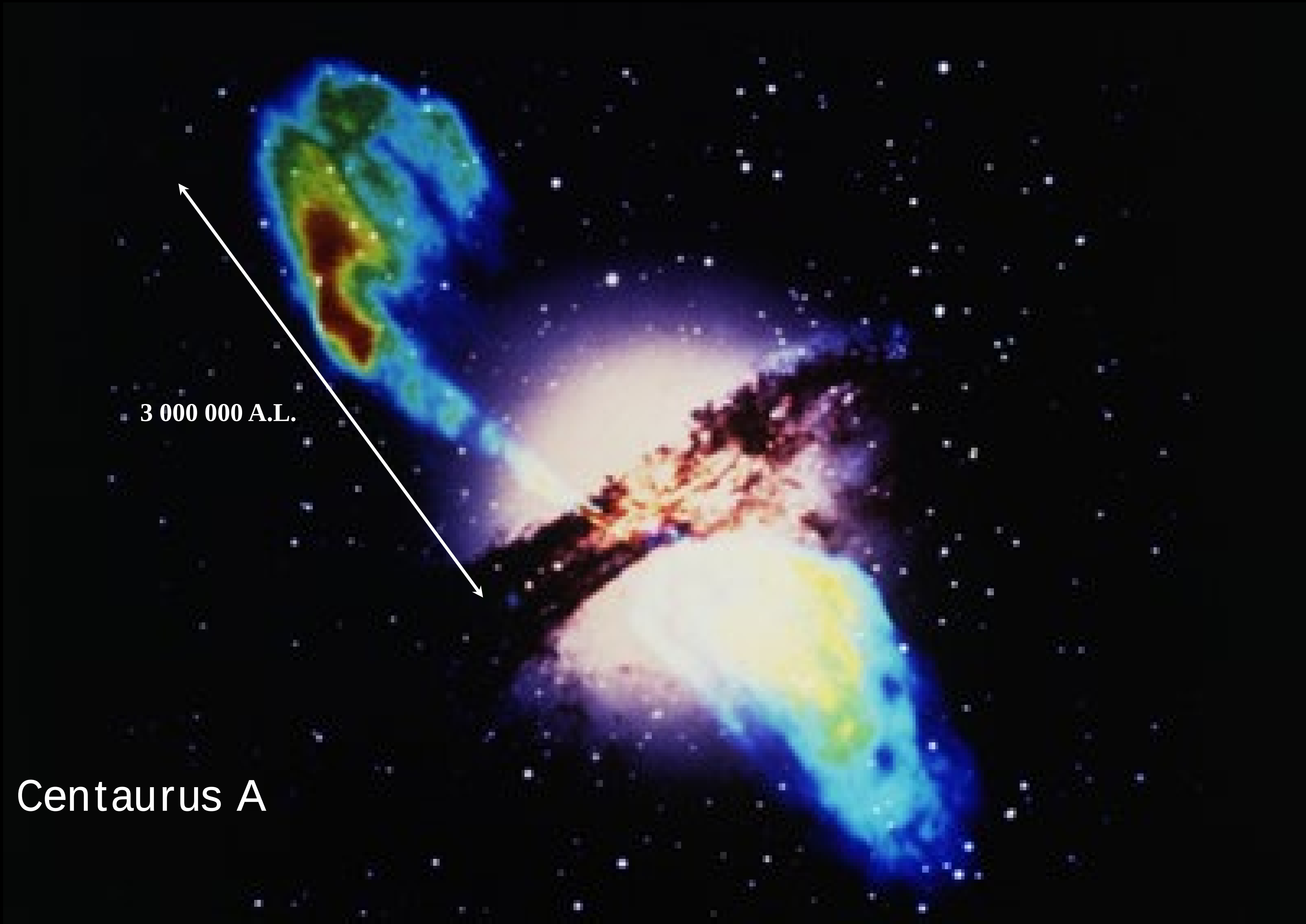
Young, hot stars in spiral arms

Andromeda (M31): Xray



Xray binaries, supernova remnants, hot gas

Radio Galaxy



Let there be light !

All what we know in astrophycics is thanks to the light !

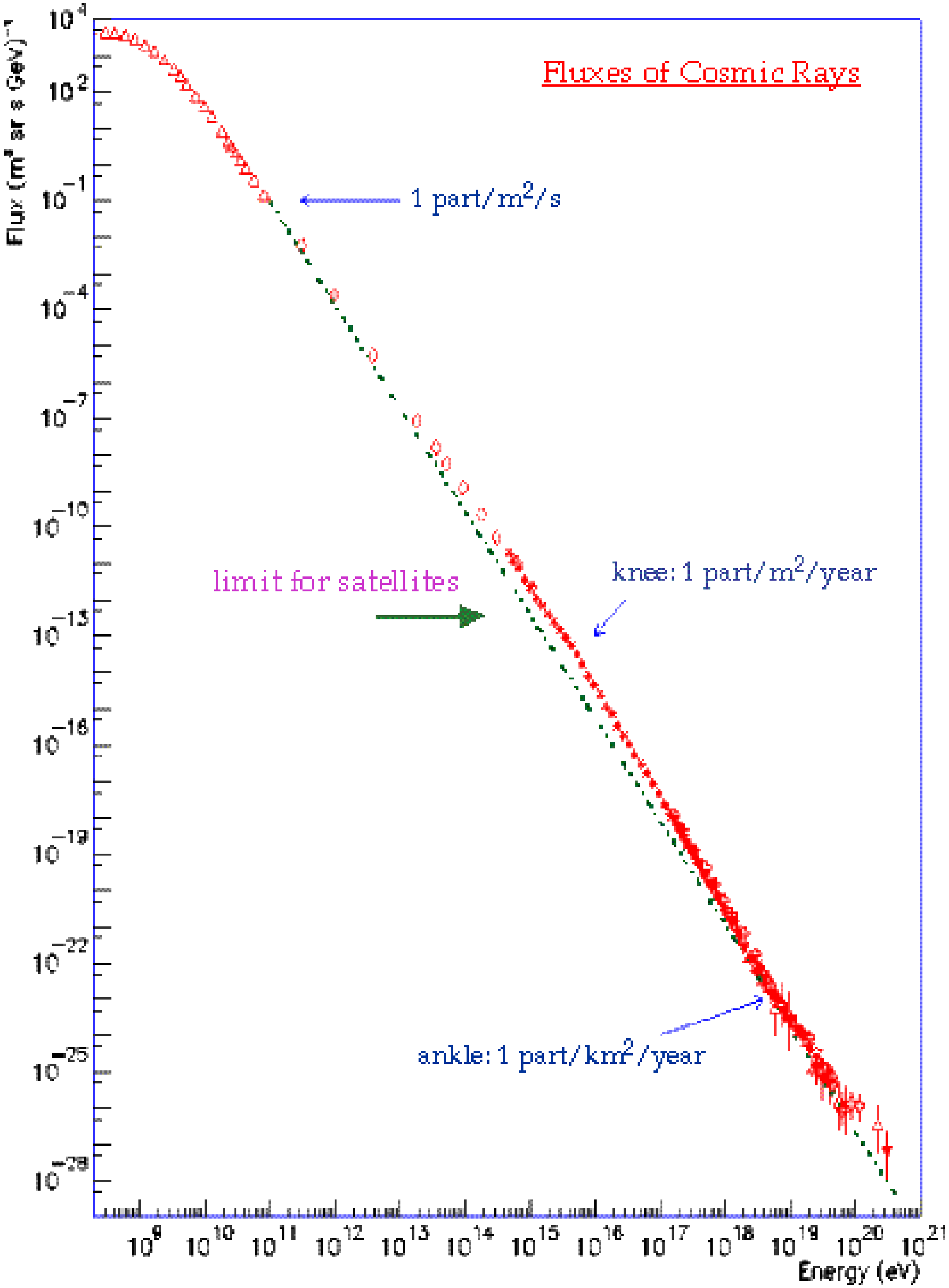
- Temperatures, stars masses, galaxies, magnetic fields, chemical composition, age of stars and structures…
- Nuclear reactions, galactic and extragalactic hydrodynamics, MHD, explosions, nucleosynthesis,
  past, future… EVERYTHING !

- Well, almost everything…
  - Non-luminous messengers : CR !
  - Rare but precious : ~ **4 CR/cm²/s**
    ~30 µg/s on entire earth (**1kg** per year !)

- CR astronomy is impossible...
  - Directions randomized by magnetic fields
  - What we would know if it was the same for photons !

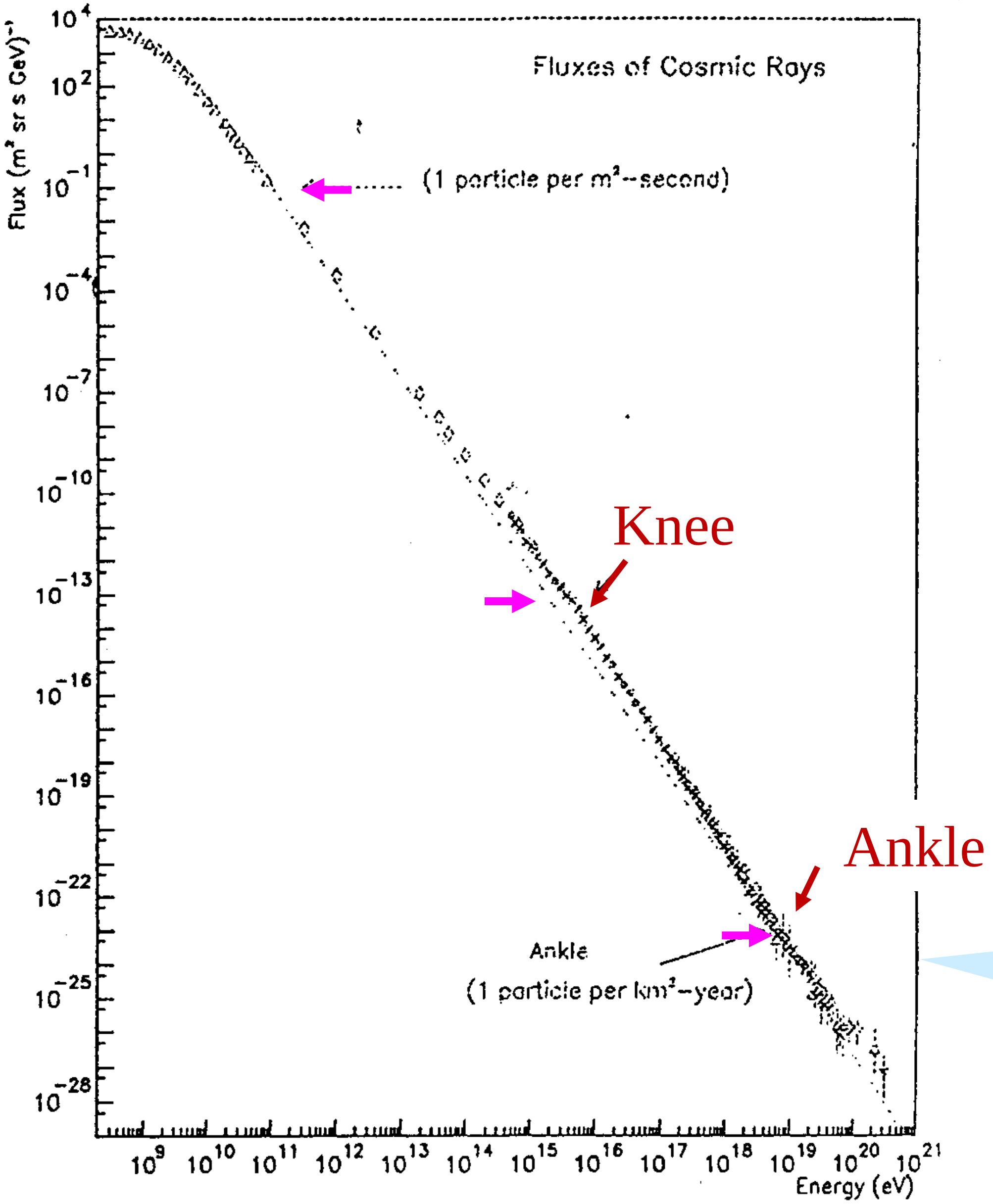
- …but not astrophysics !
  - Energy spectra and chemical composition tells us a lot…

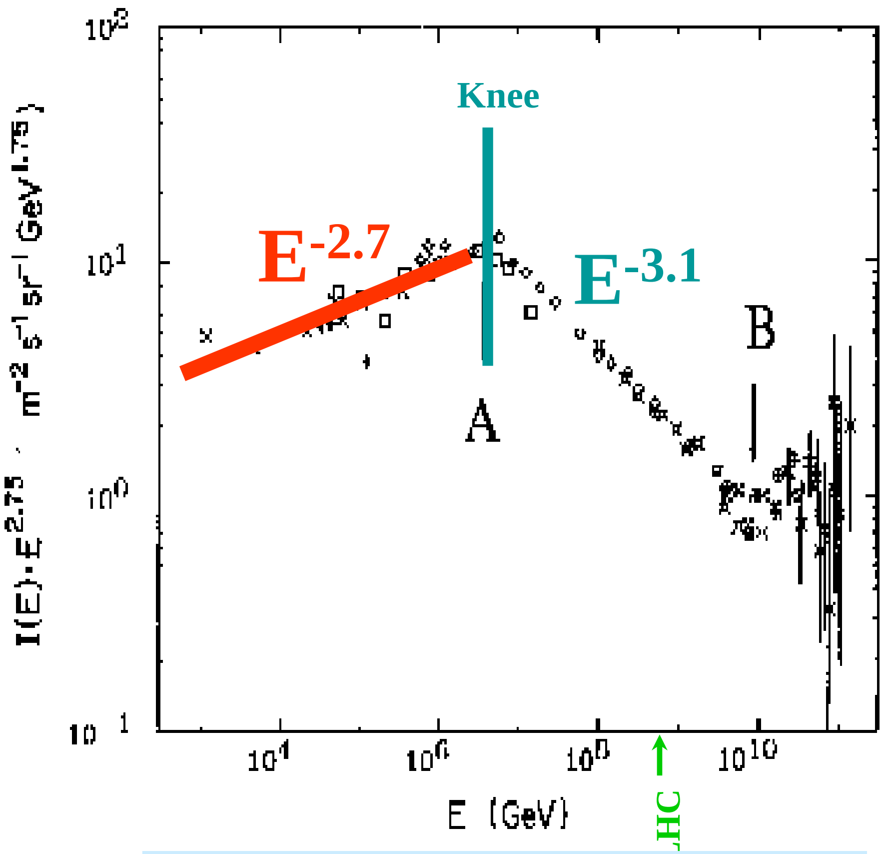
The "all particles" spectrum



- Regular spectrum over 12 decades in energy, and 32 decades in flux !!!
- Small break near 3´ 10^{15}eV : the "knee"
- An other one near 10^{18}eV : the "ankle"
- Spectrum badly known at the two extremities
  - Geomagnetic "shield" + Solar modulation
  - Extreme rareness…

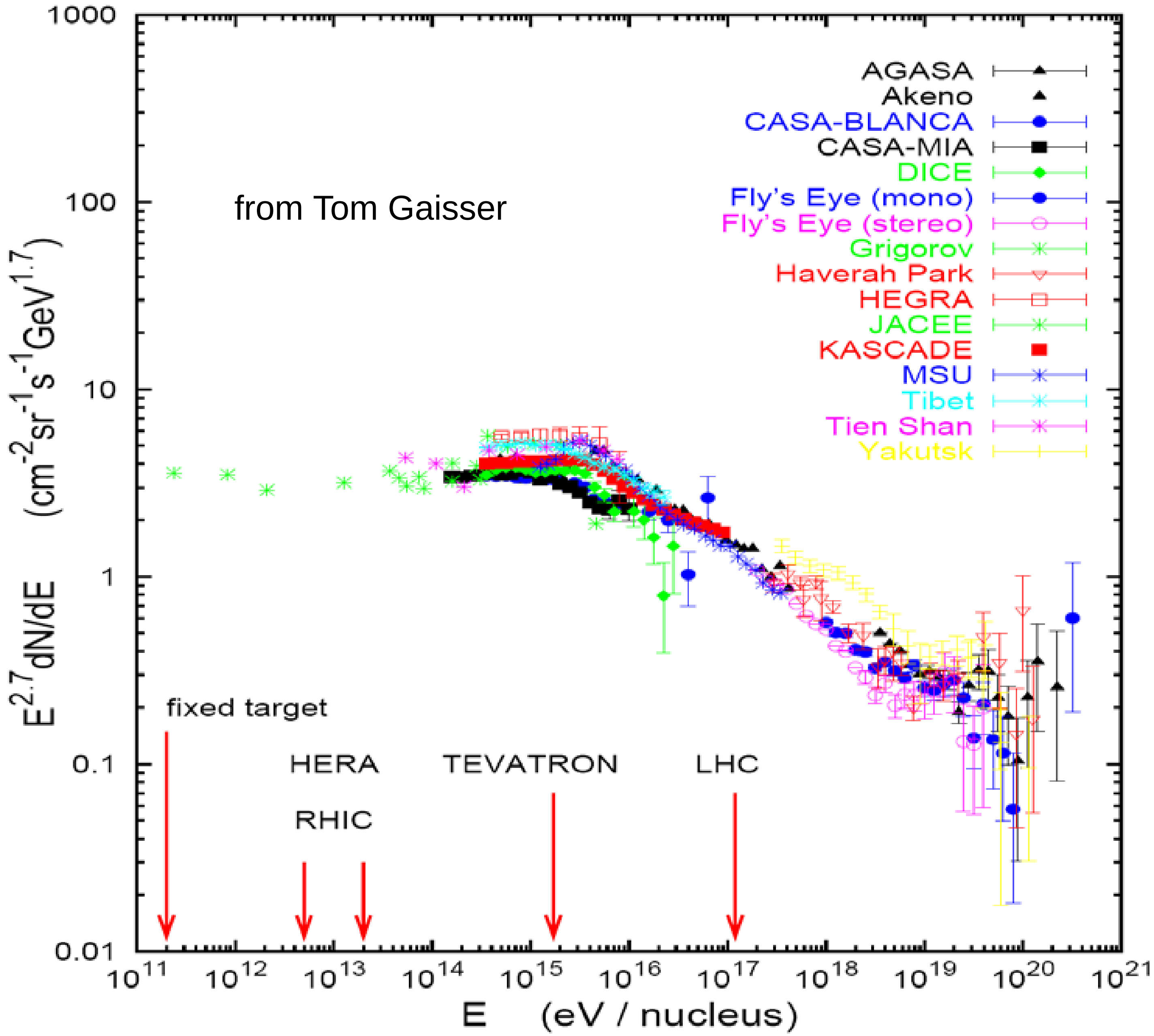
Cosmic-rays spectrum



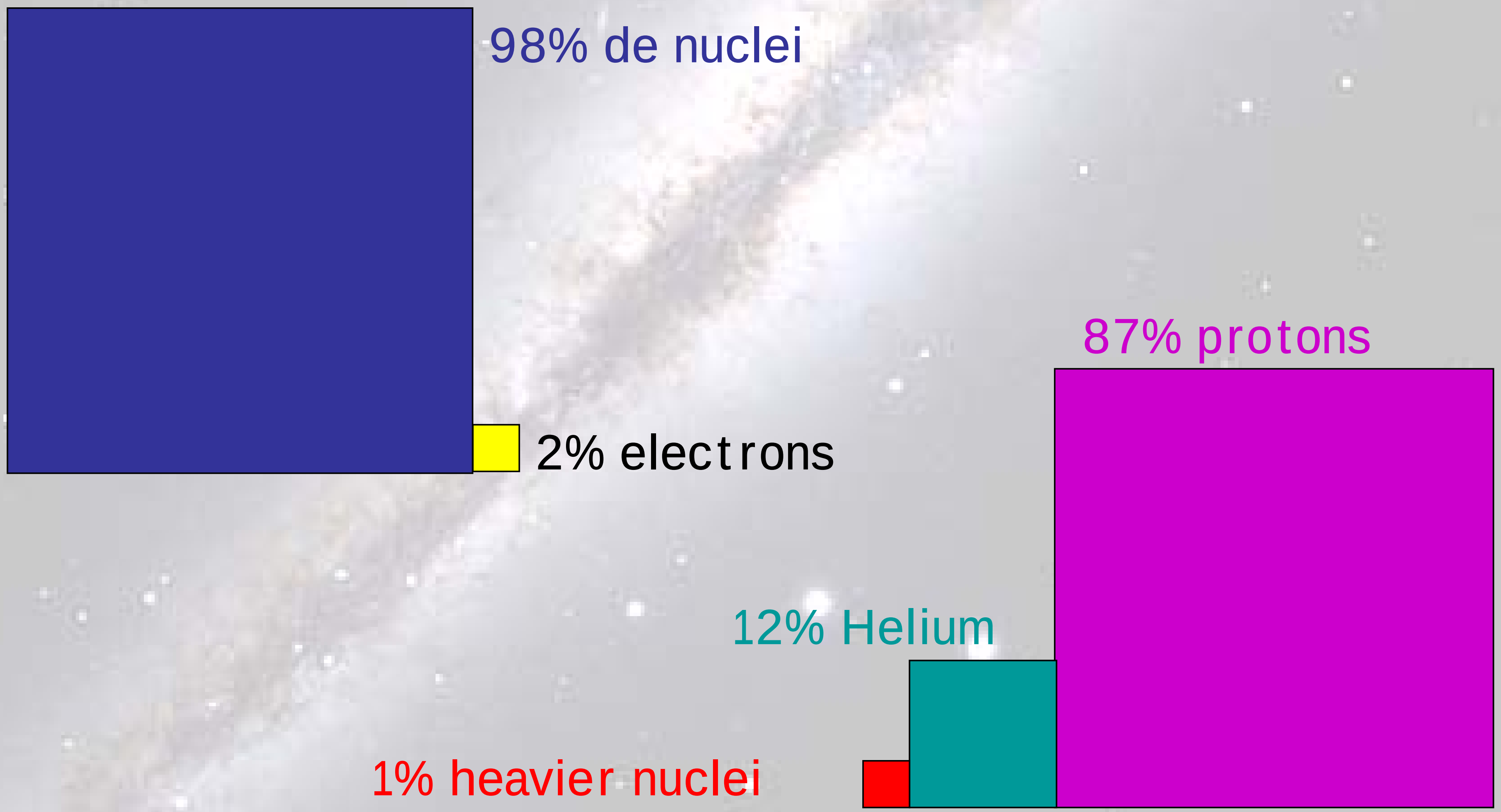


Origin ?
Galactic ?
Cosmologic ?
Astrophysics ?
New physics ?

CR Spectrum above a TeV

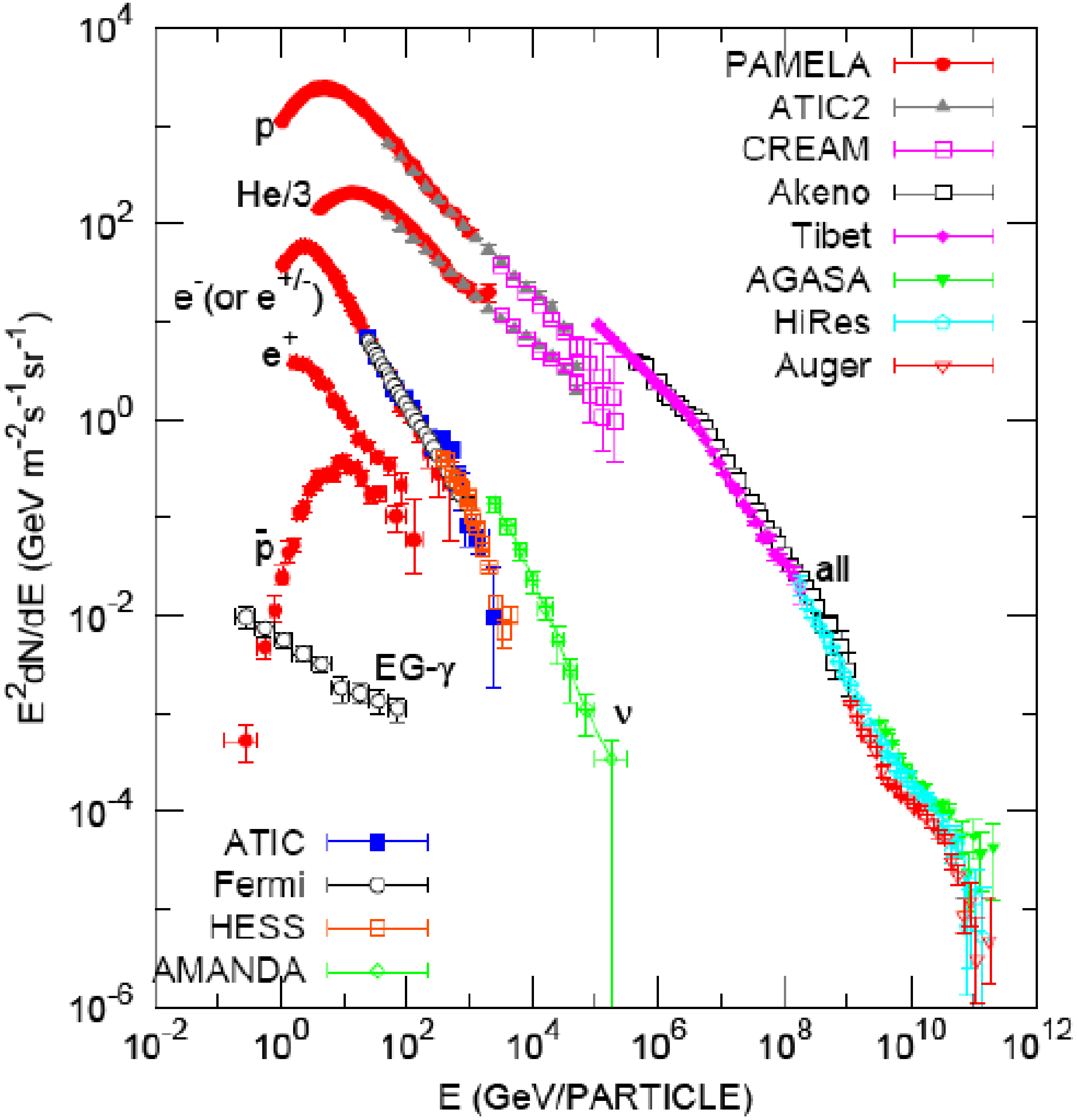


Charge cosmic rays composition



Flux : 4 RC/cm^2/s Þ 1 kg/year << 40 000 ton/year (meteorites)

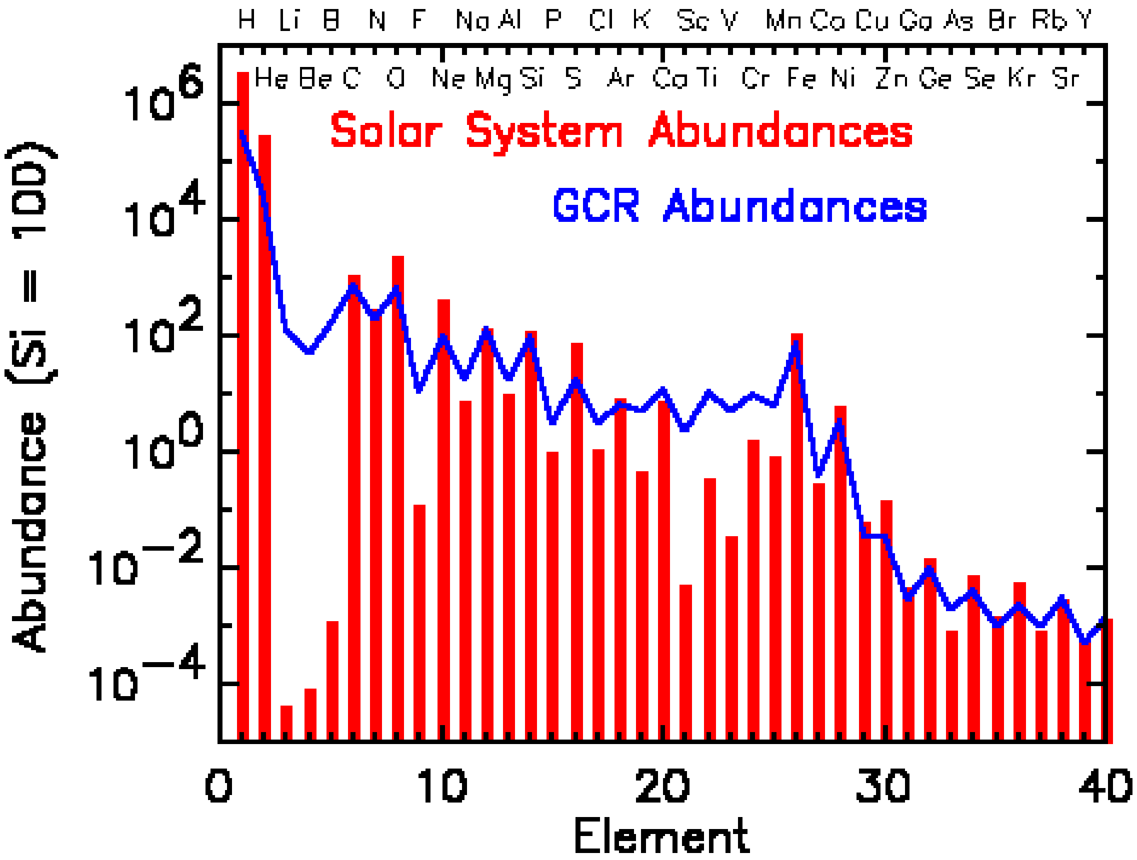
Identified spectra

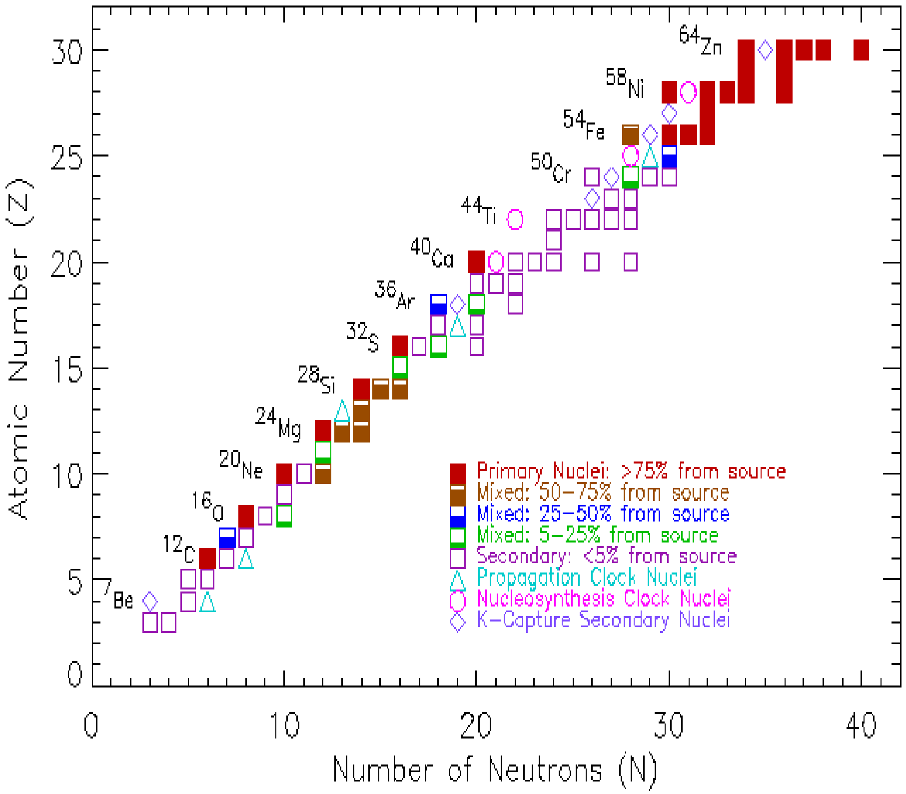


Overview of CR data on composition

- Chemical composition
  - Nuclei = 98% (H = 87%, He = 12%, "metals" = 1%)
  - Electrons = 2%
  - More or less standard compositon (i.e. solar system) except for fewer H and He, presence of secondary nuclei, and a few "anomalies"...
- Secondary atoms
  - Li, Be, B : spallation of C, N, O (+ nuclei below the Fe peak)
  - Nuclear thicknesses traversed by CR : $X_{CR} = 6$ to 10 g/cm^2
- Isotopic anomalies
  - $^{22}Ne \rightarrow$ link with massive stars
- Cosmic clocks
  - $^{10}Be \rightarrow ^{10}B$, $\tau \approx 4 \times 10^6$ years (as well as ^{26}Al, ^{36}Cl, ^{53}Mn, ^{54}Mn, ^{59}Ni)
  - $\tau_{RC} \approx 2 \times 10^7$ years
  - $\frac{X_{RC}}{c\tau_{RC}} \approx 0.2$part/cm$^3 \Rightarrow$ CR halo extention ($\approx$ 3-7 kpc)

Nature of cosmic rays

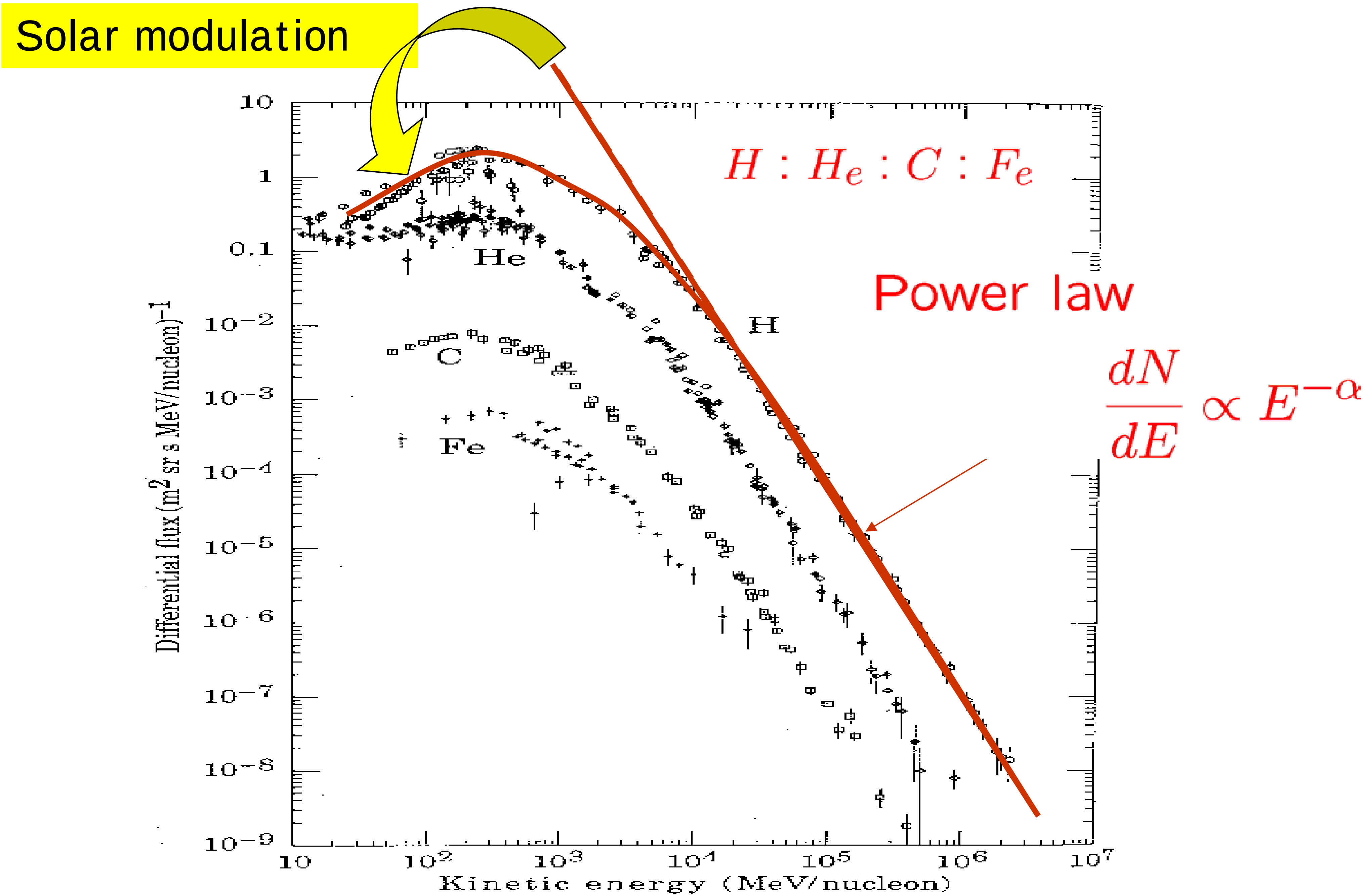




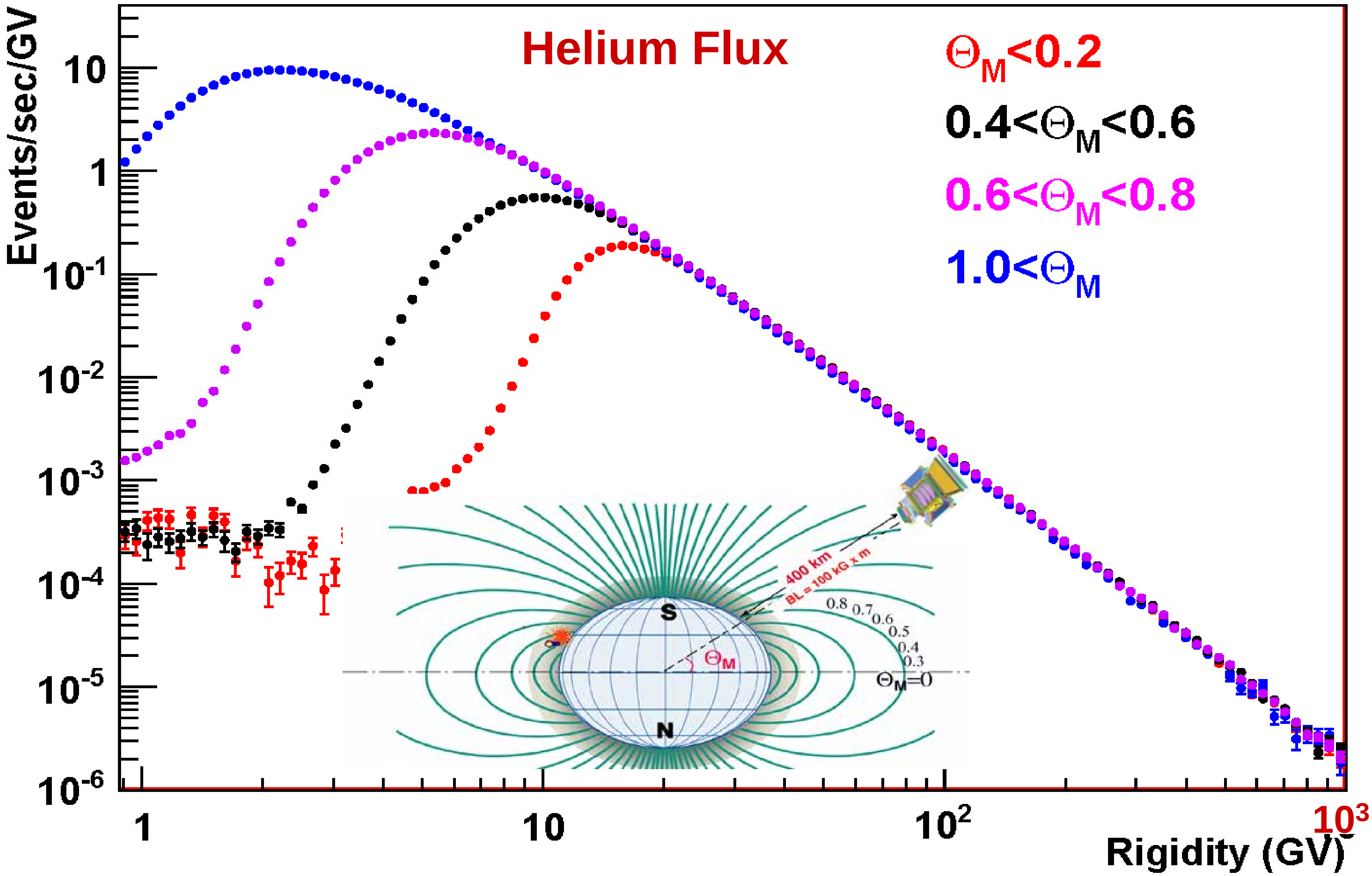
Abundances different in CR and local measurements
(Li Be B and Sub-Fe)

CR undergo spallations and produce secondary CR

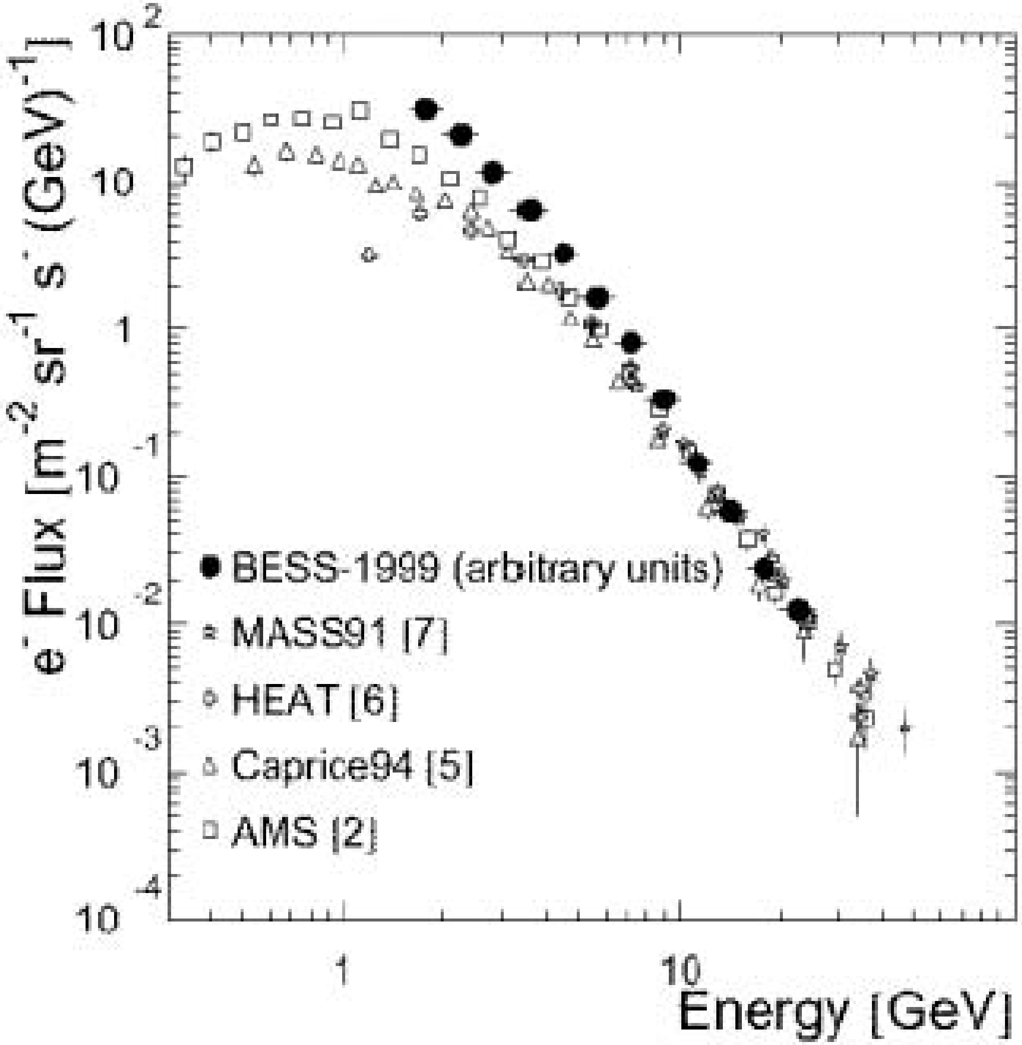
Nature of primary cosmic rays

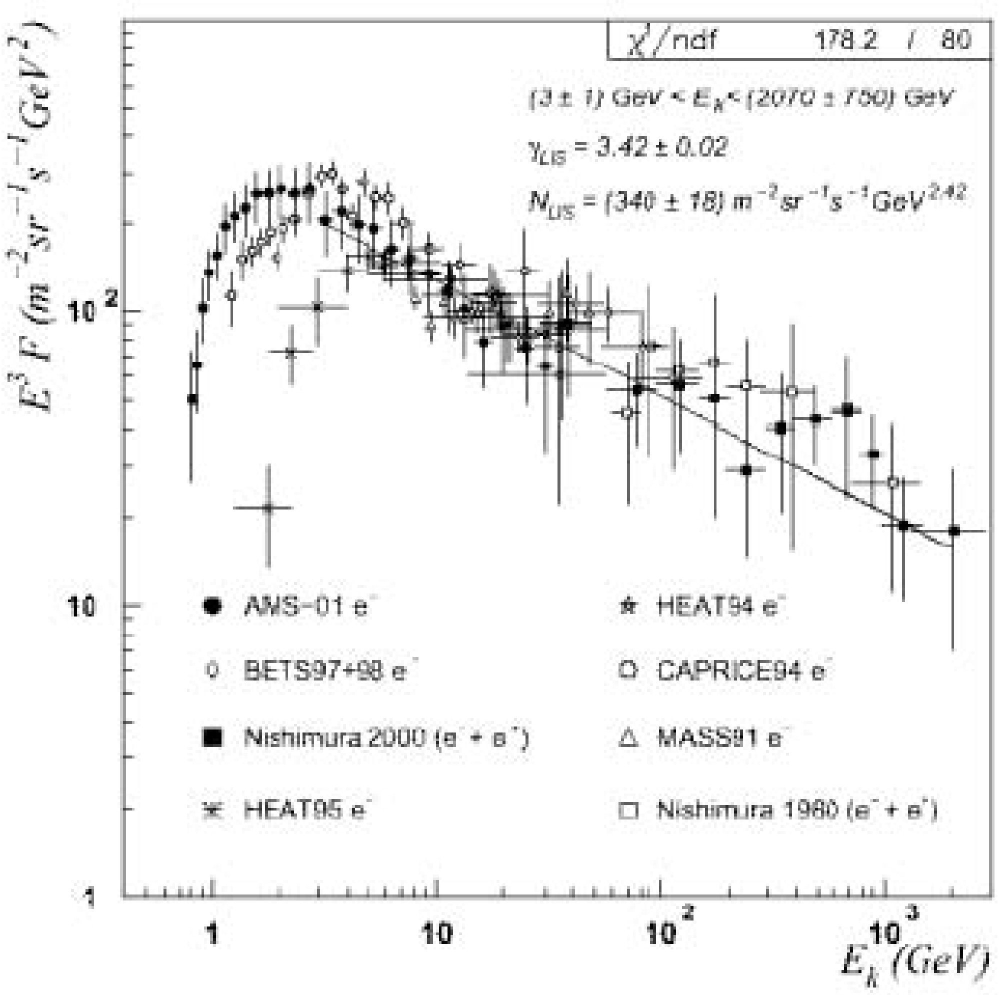


Data from AMS on ISS



Electron primary flux



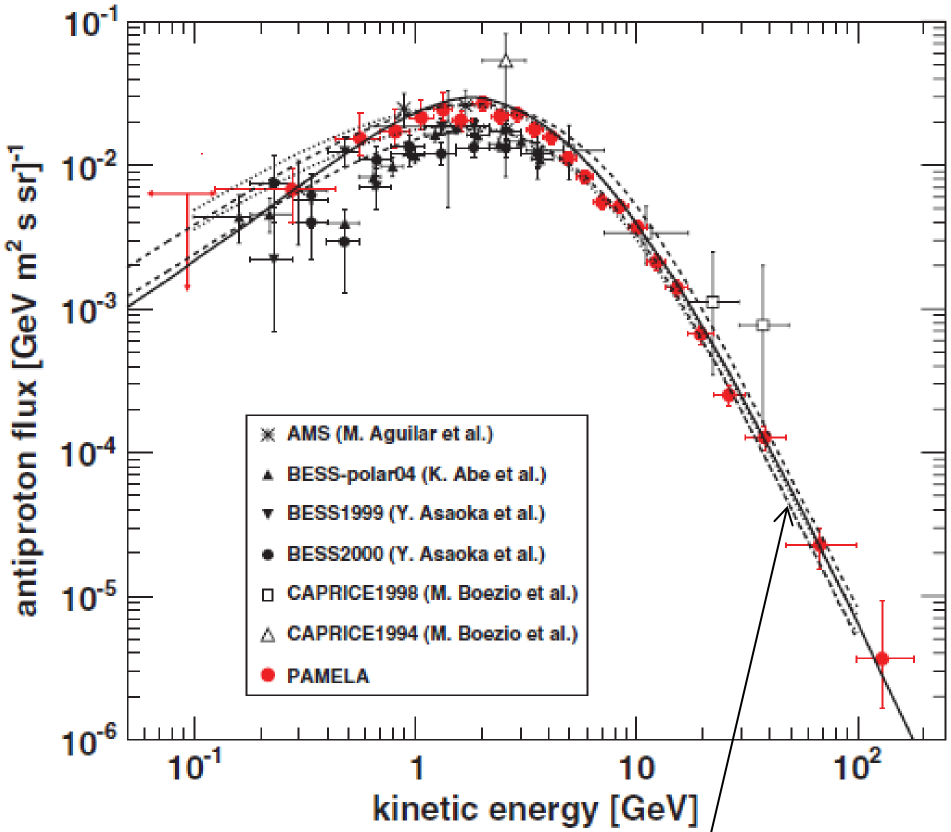


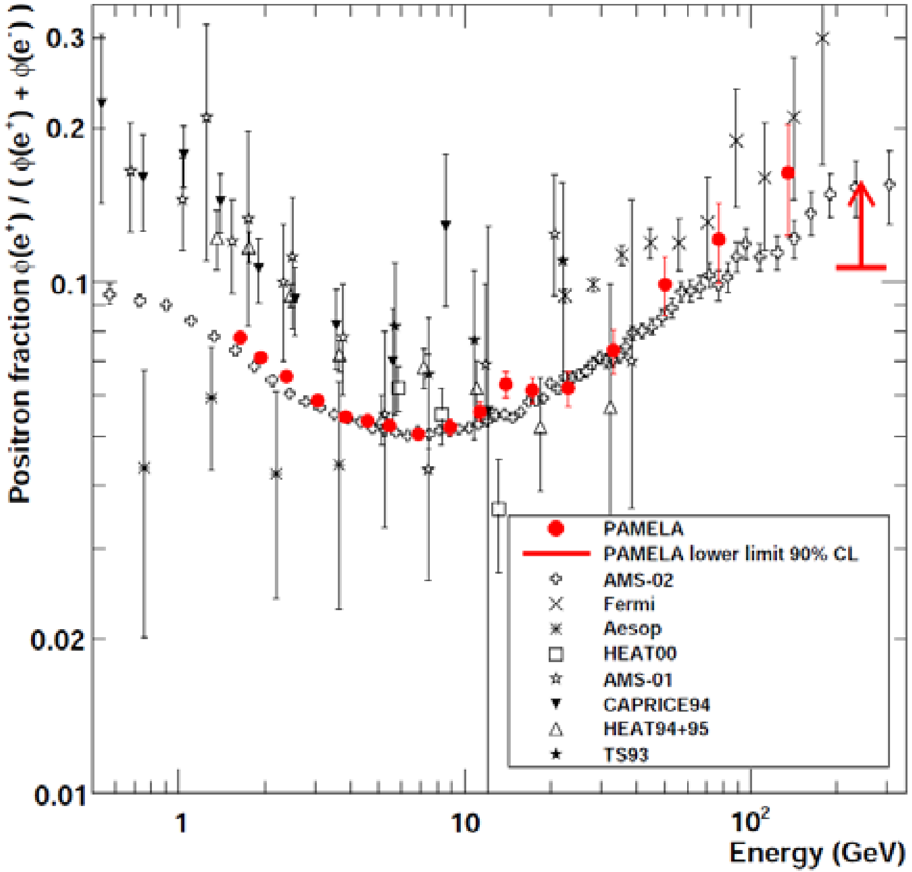
Nature of primary cosmic rays

Antimatter ?

$$\bar{p}$$

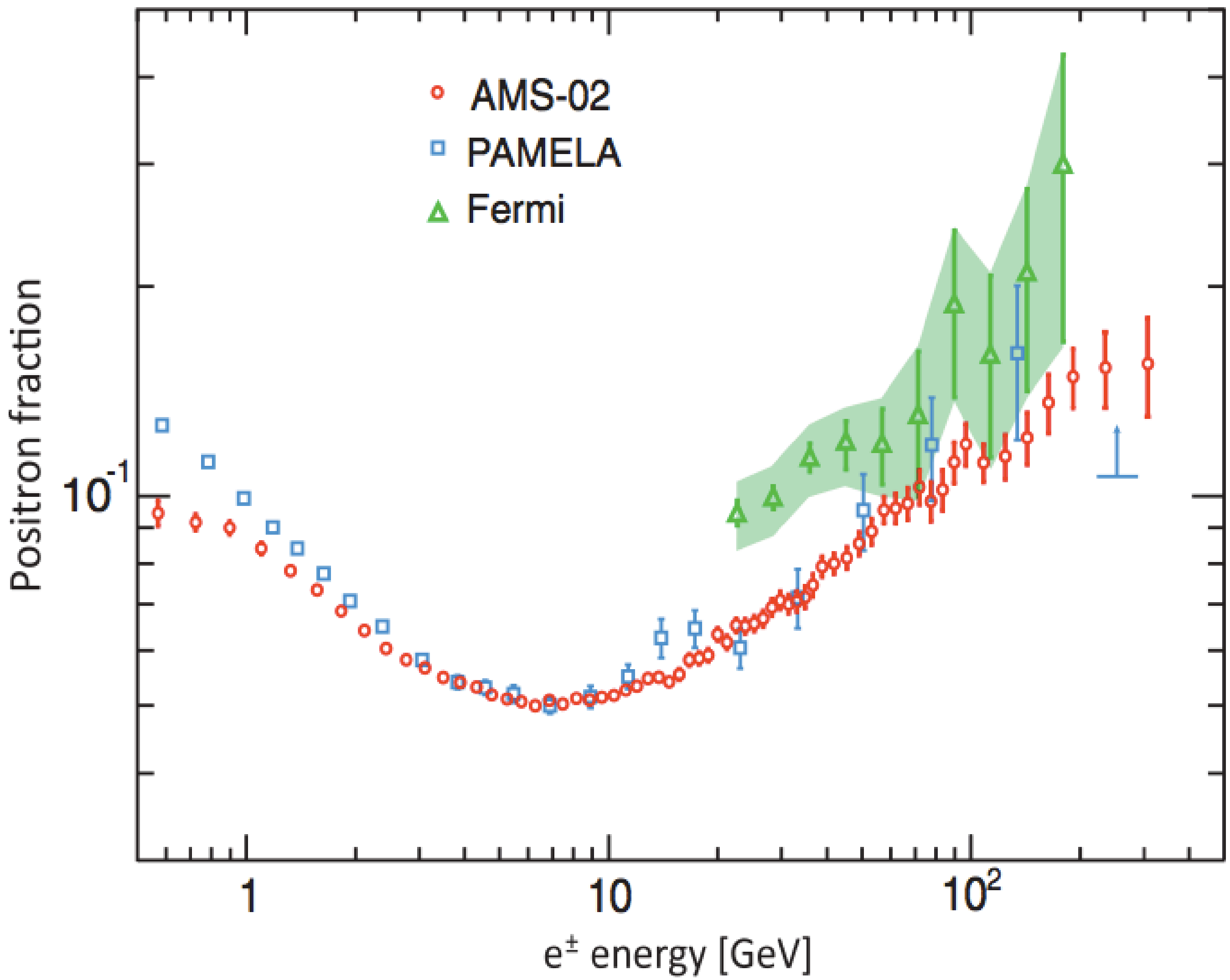
$$\frac{e^+}{e^+ + e^-}$$



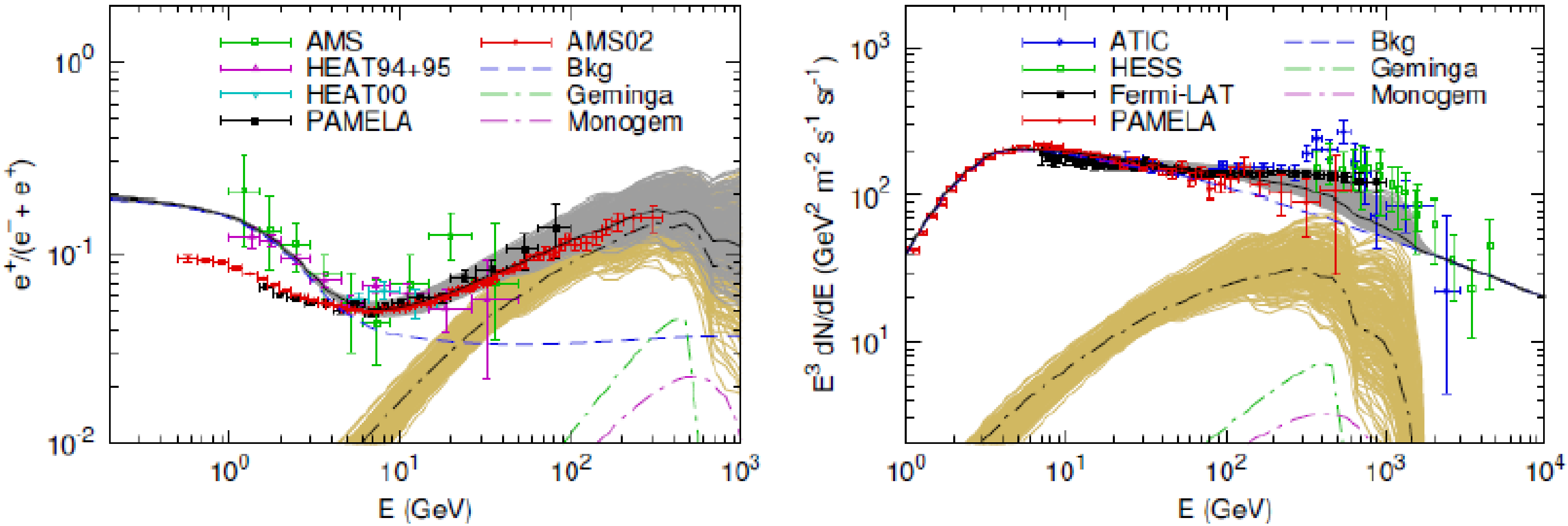


Pure secondary (standard) propagation

Positron fraction in primary flux



Antimatter search (Dark Matter ?)



… or rather a boring "local" pulsar spoiling physicists hopes !

Gamma rays

- Gamma-rays observed ® TeV
- Spectrum ± understood un to MeV.
- Above, the diffuse spectrum and that of sources are very "hard", in $1/E^2$ and of still unclear origin...

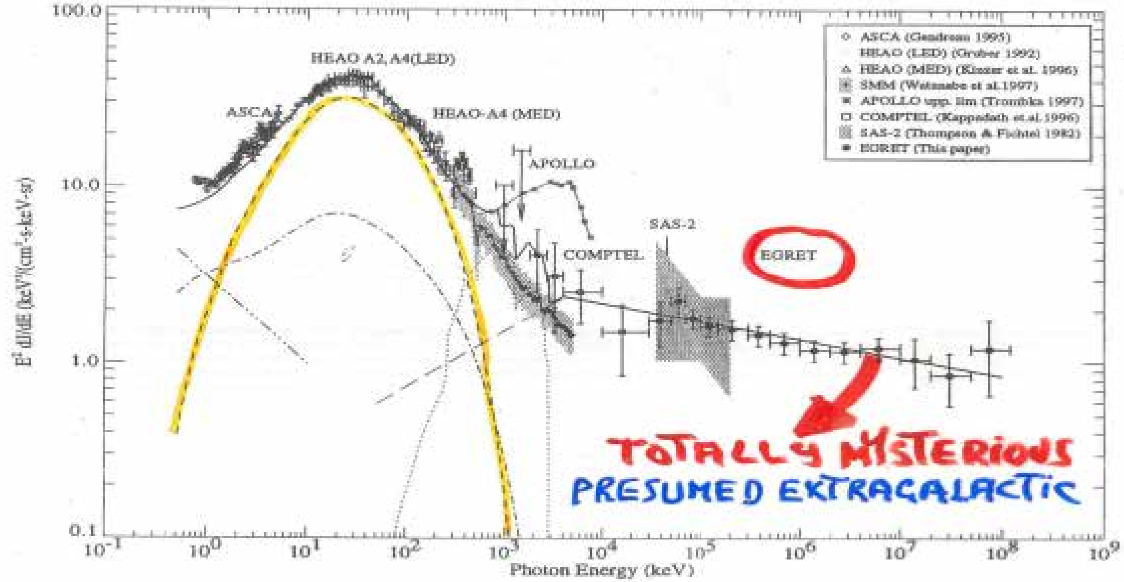
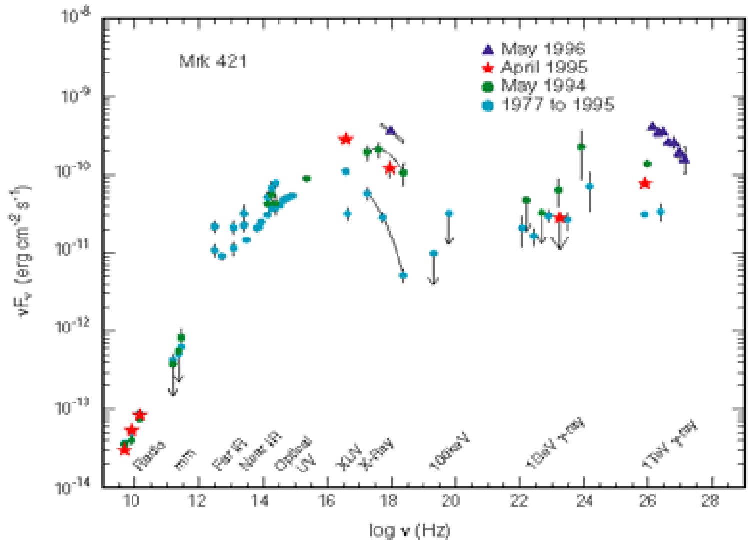


FIG. 10.—Multiwavelength spectrum of the extragalactic gamma-ray spectrum from X-rays to high-energy gamma rays. The estimated contribution from Seyfert I (*dot-dashed line*), and Seyfert II (*dashed*) are from the model of Zdziarski (1996); steep-spectrum quasar contribution (*triple-dot-dashed line*) is taken from Chen, Fabian, & Gendreau (1997); Type Ia supernovae (*dotted line*) is from The et al. (1993). The blazar contribution below 4 MeV (*long-dashed line*) is derived assuming the average blazar spectrum breaks around 4 MeV (McNaron-Brown et al. 1995) to a power law with an index of ~ −1.7. The thick solid line indicates the sum of all the components.



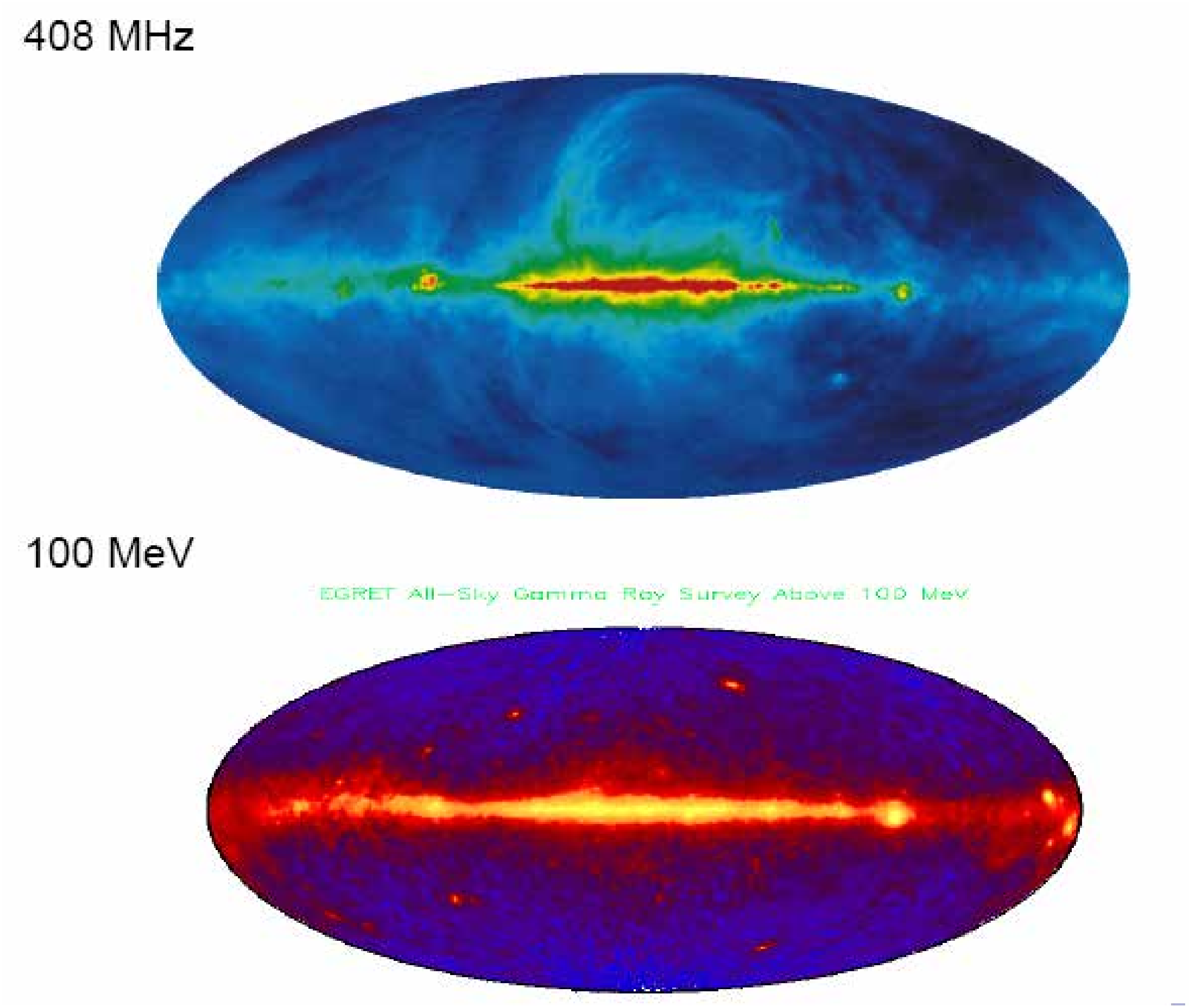
Why all this non thermal equilibrium radiation?

Gamma, diffuse emission

Emission due to:

- the interactions of cosmic electrons with:
  - the magnetic fields (synchrotron radiation dominates the radio emission of the Galaxy up to a few GHz)
  - interstellar Matter (ISM); bremsstrahlung important bellow 100 MeV
  - Interstellar photon: Inverse Compton above GeV
- the decay of p^0 produced when CR interact with protons and nuclei
  - p^0 ! gg above 100 MeV
  - Concomitant emission of n in the decay of p§

Gamma, diffuse emission



Galactic or Extragalactic CR ?

Definite answer from EGRET in 1993 !

Hypothesis: if CR are extra or metagalactic, the density of CR is identical in our Galaxy and in its satellites

- Radio observations radio give the mass of gas M_H in the SMC
- M_H implies a measurable flux for SMC of: $2.5 \times 10^{-7} cm^{-2}.s^{-1}$

$$F_\gamma \propto M_H N_{CR} R_q$$

- EGRET gives an upper limit (at 95%CL): $< 0.5 \times 10^{-7} cm^{-2}.s^{-1}$
- The CR density is 5 times smaller within SMC

Cosmic rays are indeed Galactic !

The general problematic

- Thermal speeds ! RCUHE (few 10^{20} eV)

Produce them

- From top to bottom (decay...)
- From bottom to top (acceleration)

Preserve them

- Energy losses (Synch., IC, p, pairs...)
- Destruction (photo-dissociation...)
- Escape probabilities

Propagate them

- Propagation in ISM and IGM (mag fields: deflection, confinement...)
- Re-acceleration

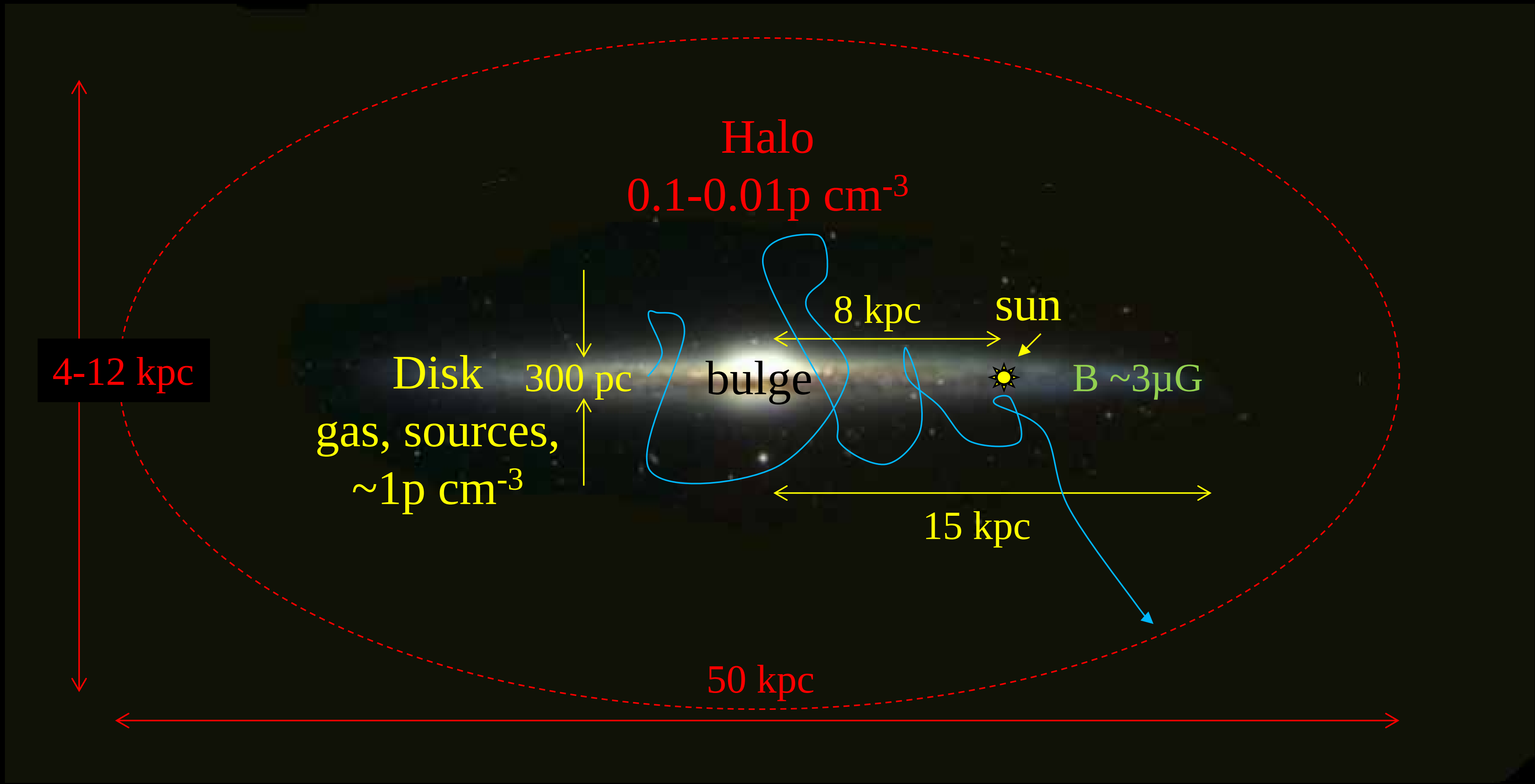
Detect them

- Balloons, satellites...
- Air showers...
  - Cherenkov telescopes
  - Surface & Fluorescence Detectors

Propagation medium, IGM, ISM and atmosphere

Dimensions of the Milky Way

1 pc ¼ 3 l.y. ¼ 3 £ 10^{16} m

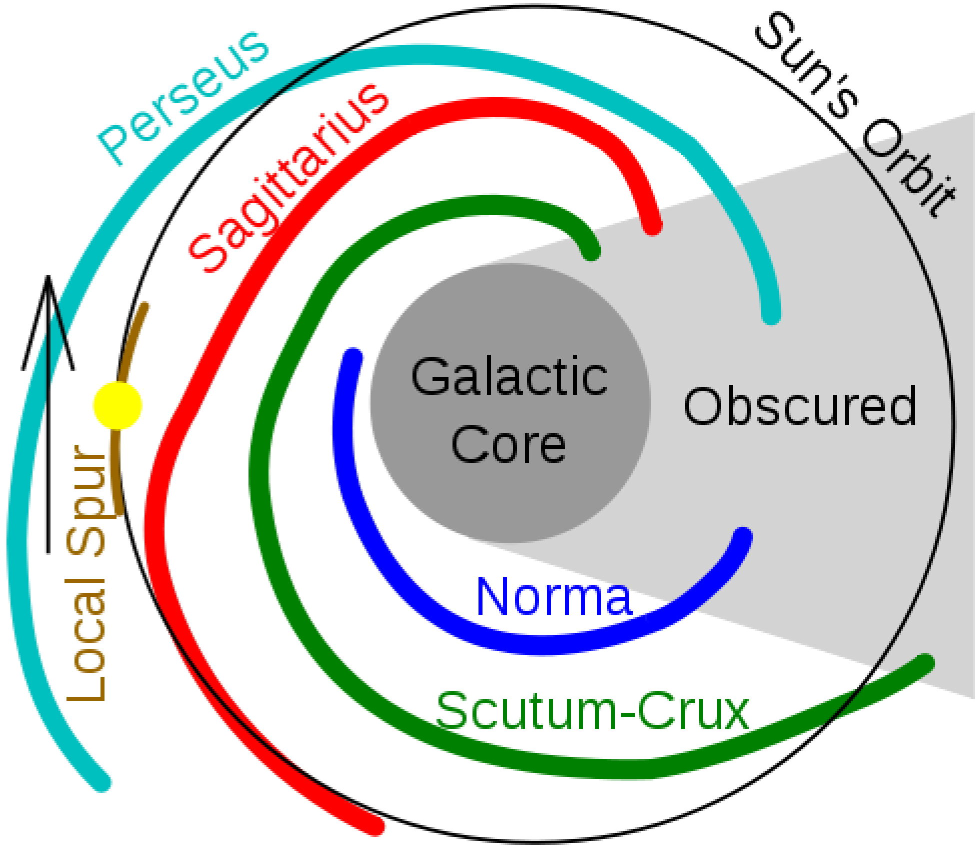


Milky Way, a spiral galaxy

Milky Way, a spiral galaxy

Local spur and
neighboring arms
) local matter and
B field inhomogeneity.

Mean "regular" B field
~ 3µG roughly parallel
to spiral arms, more
intense in between arms.



Milky Way, a spiral galaxy

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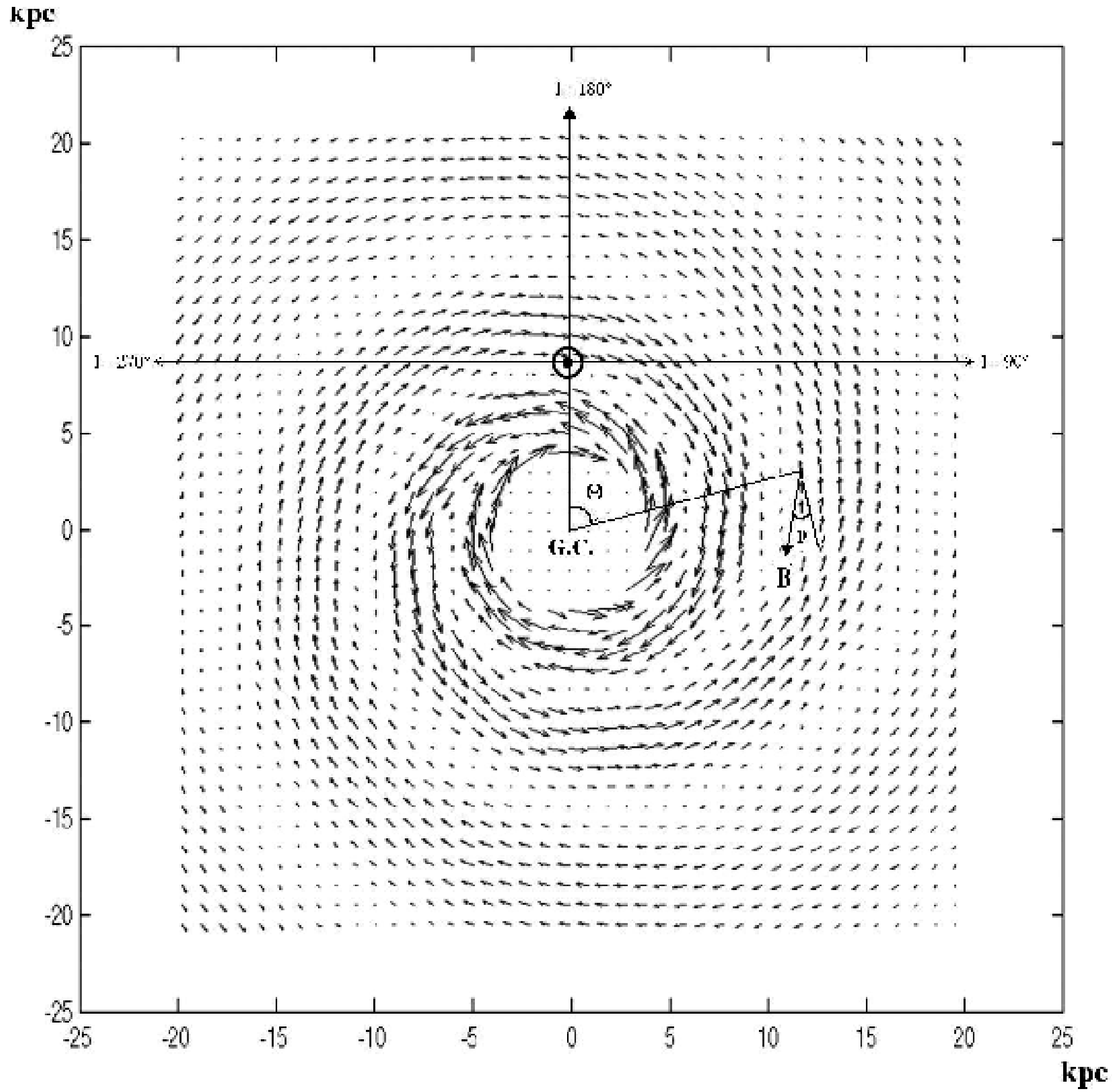
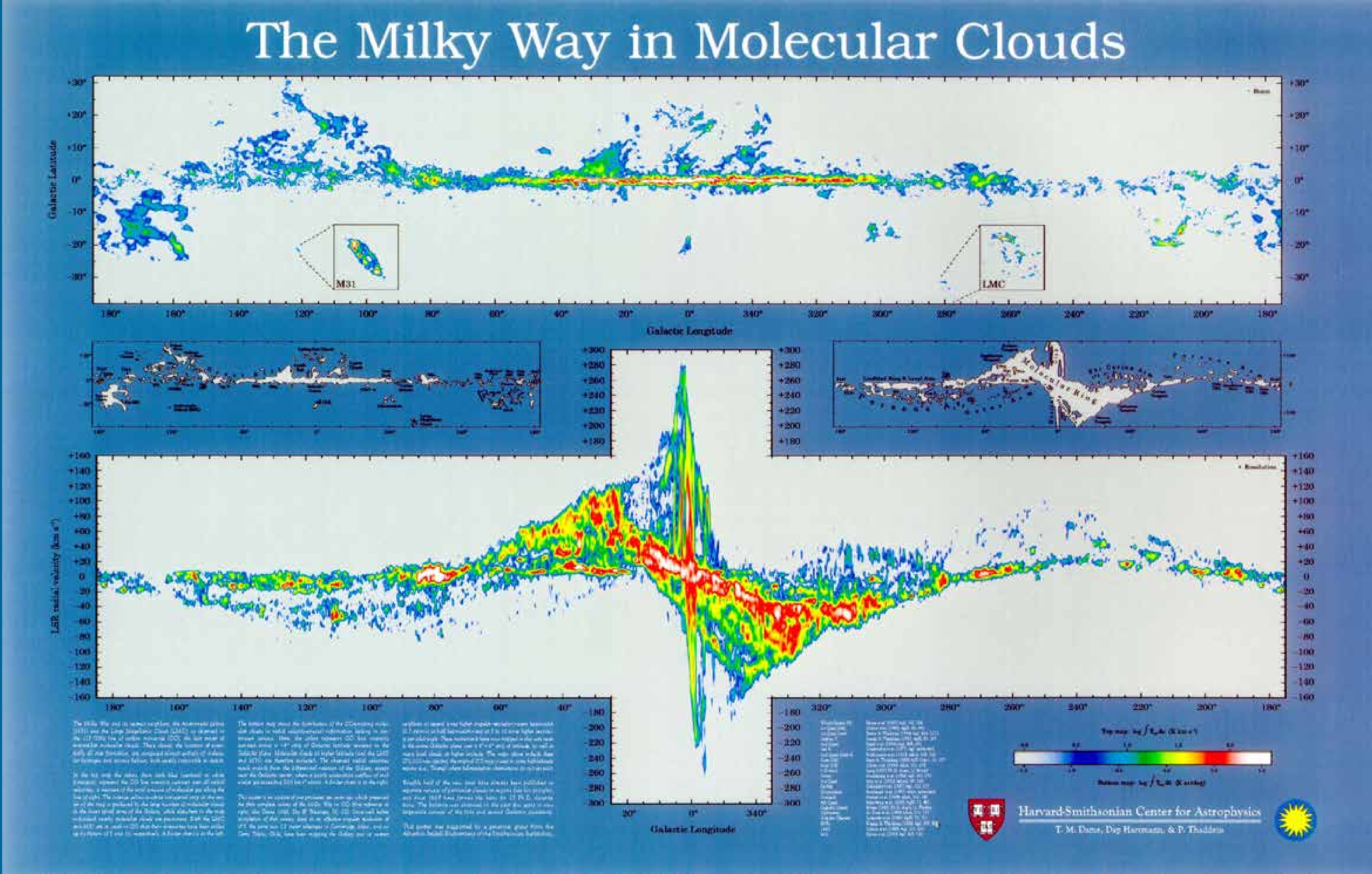
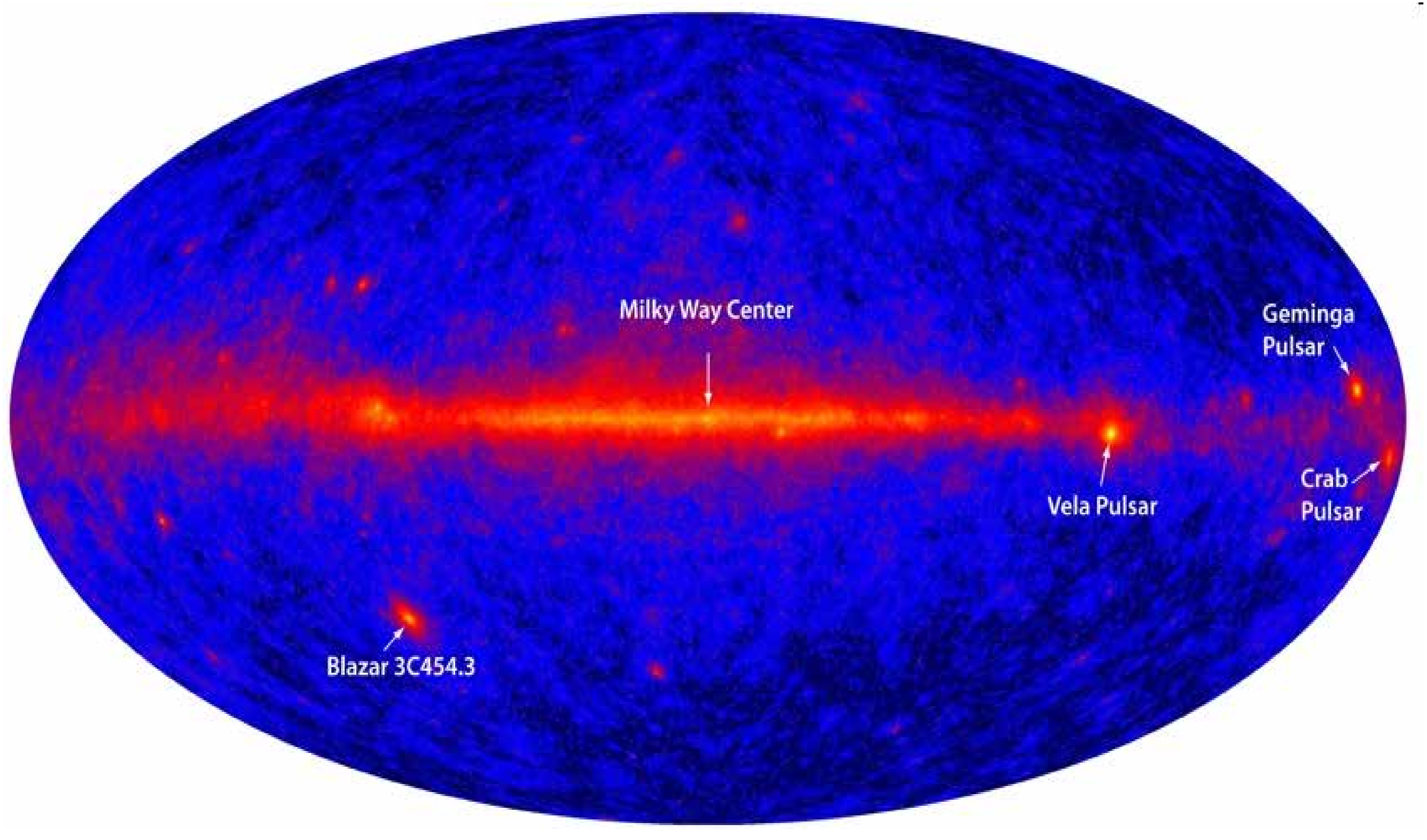


A thick target…

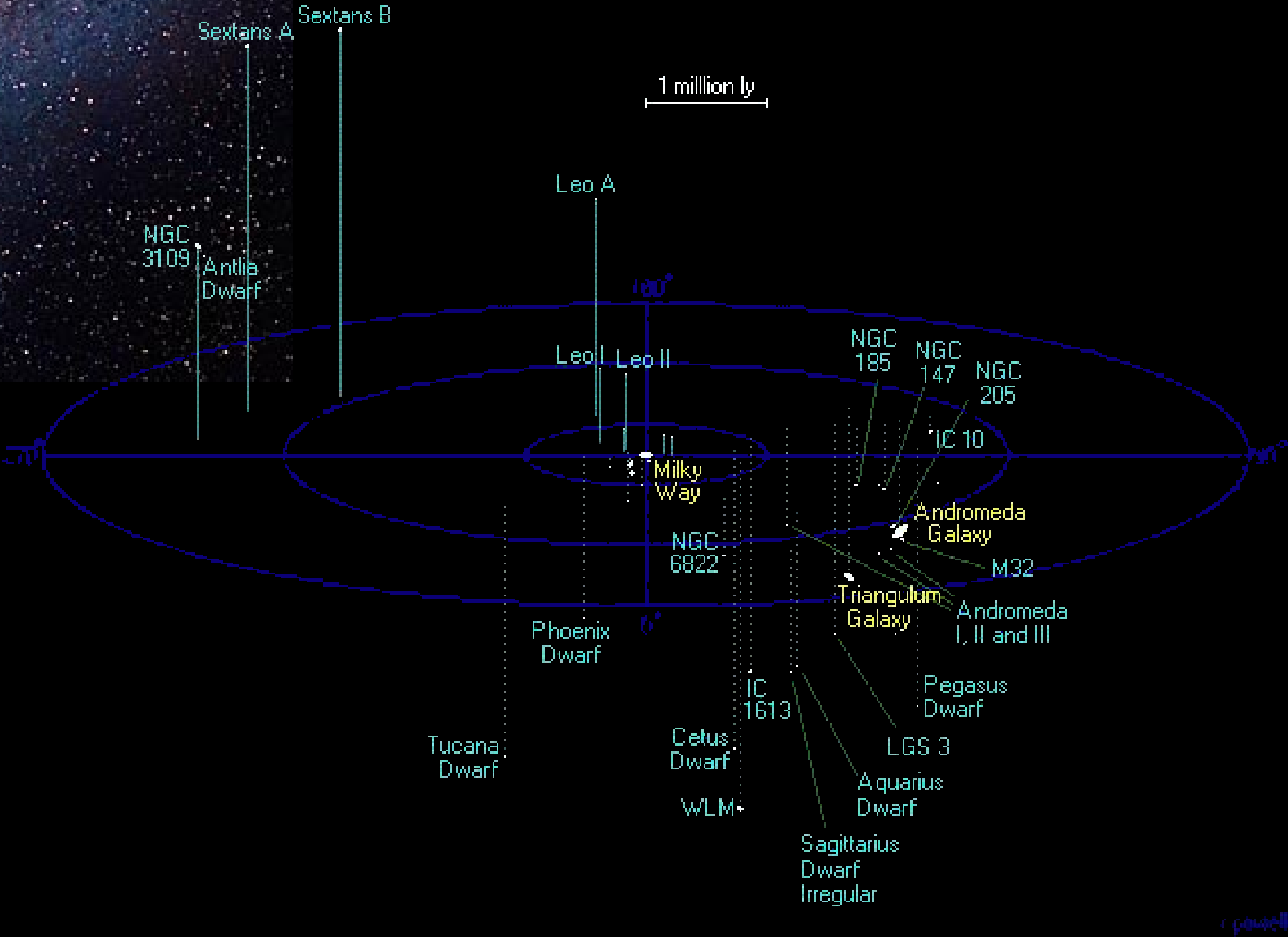


A thick target

- Diffuse gamma-ray emission from galactic CR interaction with matter (mostly molecular H clouds).



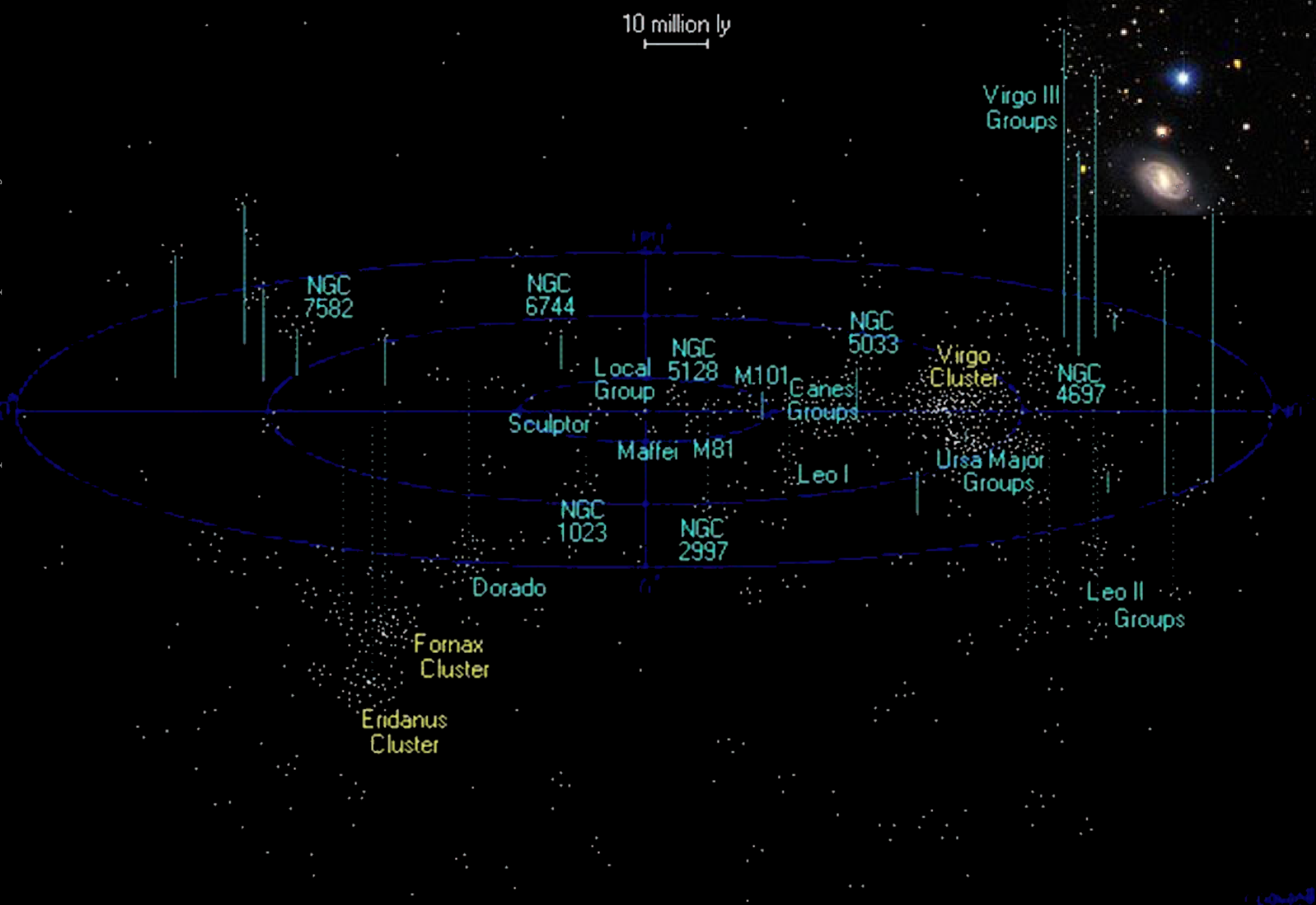
The nearby islands...



Andromeda (M31)
A twin of our Milky Way slightly larger and (only) distant by 780kpc.
Many small (dwarf) galaxies are orbiting around these twins.

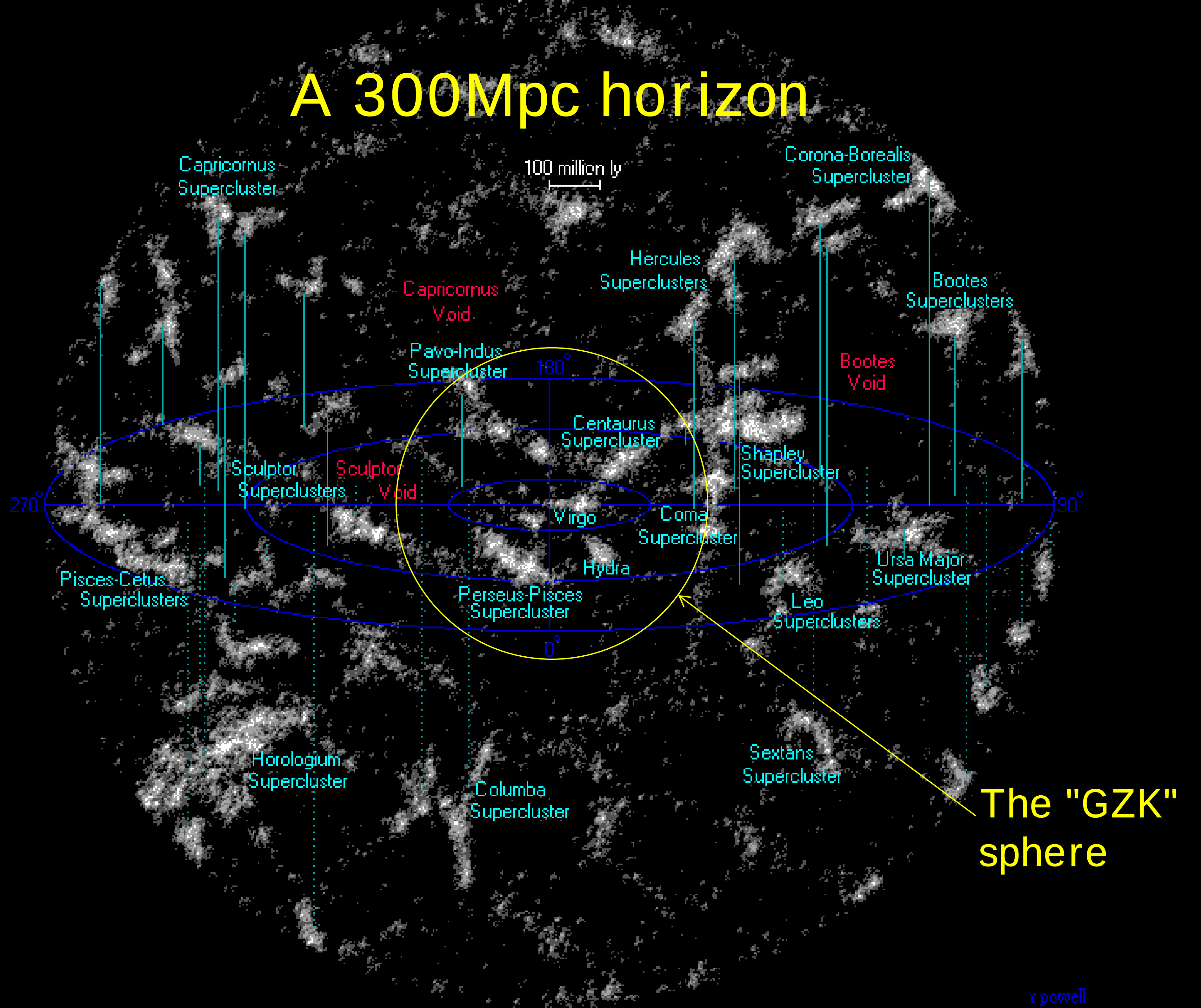
The local group and the Virgo cluster

Our local group is at the periphery of the large Virgo supercluster (~2000 galaxies) at ~20Mpc

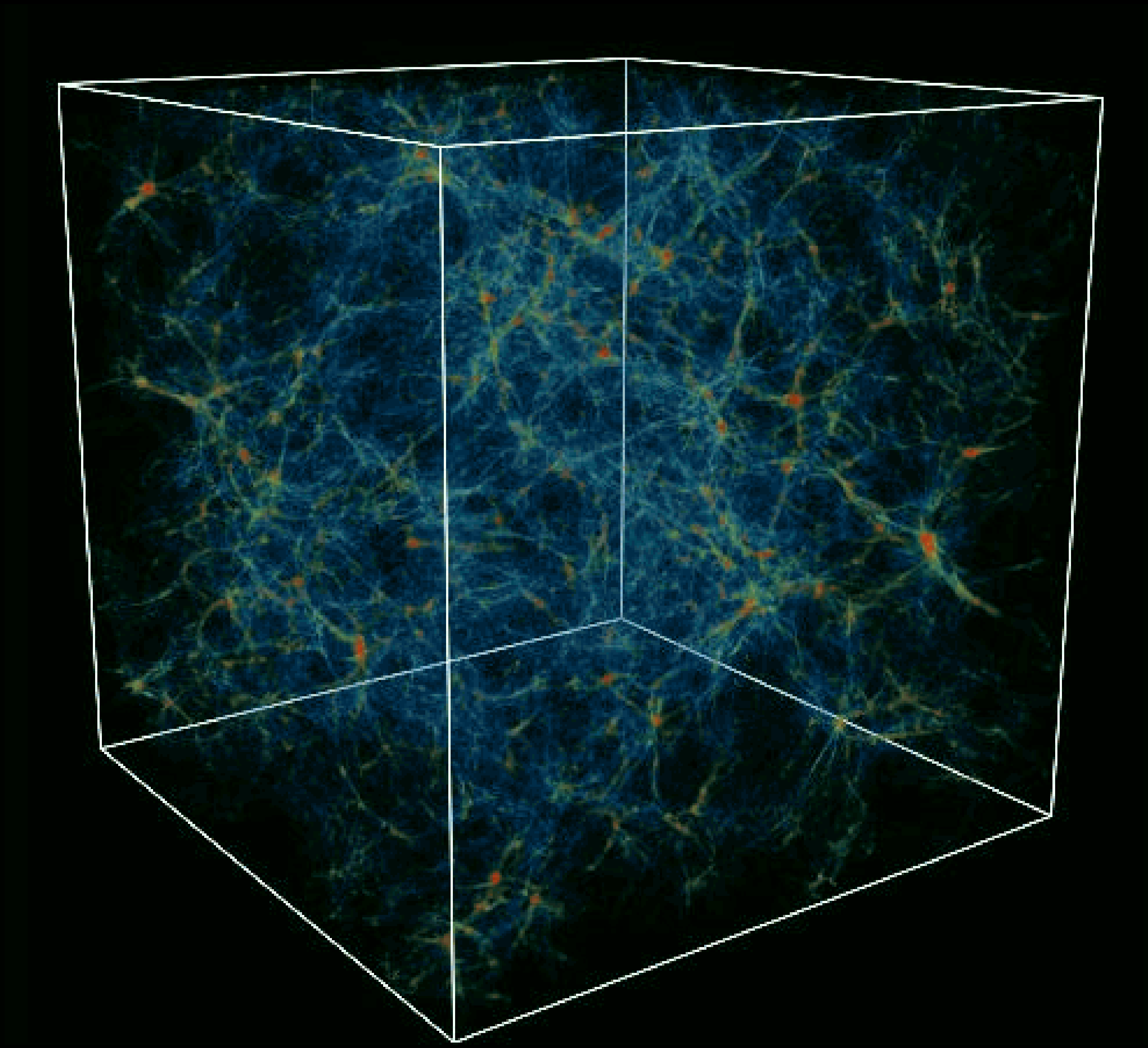


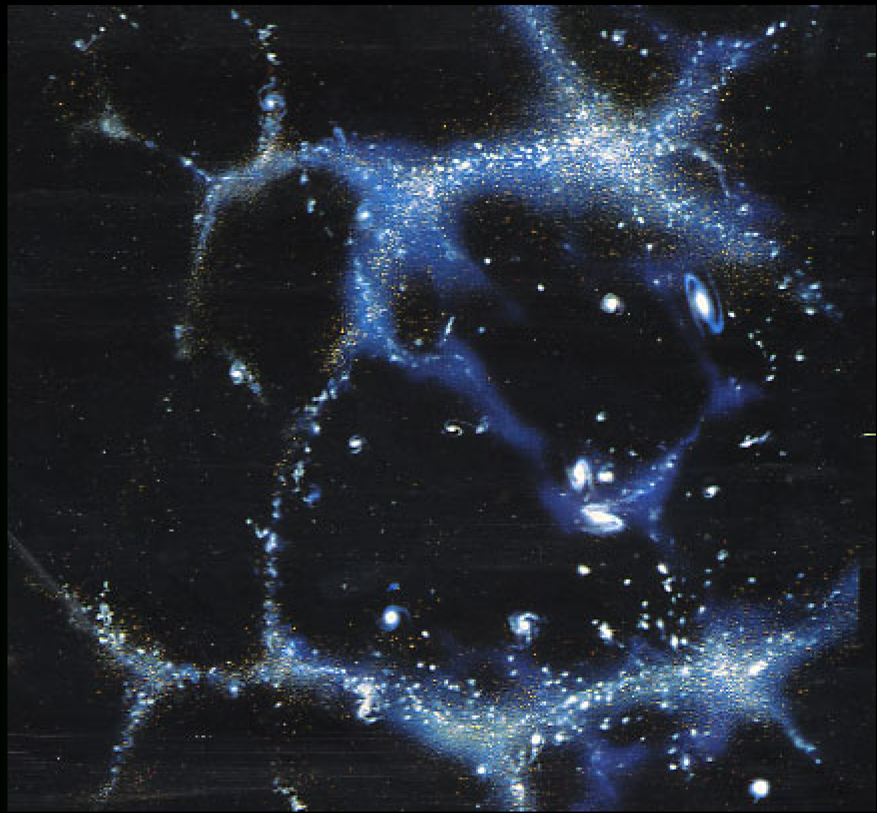
Another super cluster: Abel 1689

A 300Mpc horizon



Large scale filamentary structures

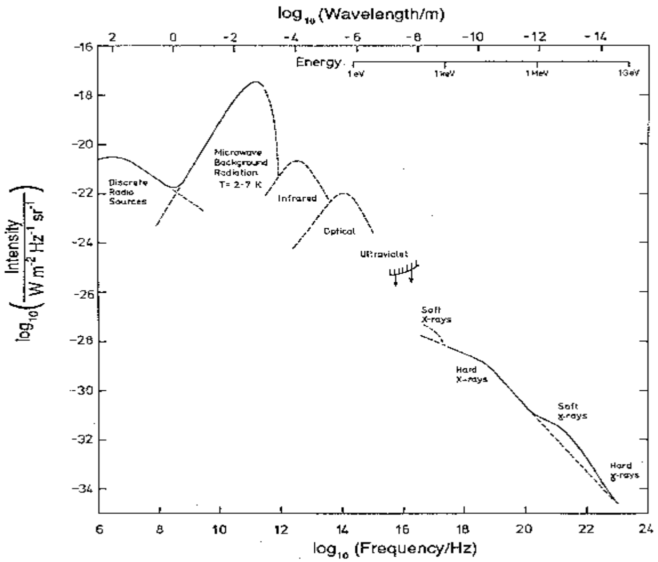




Vacuum is not emptiness !

- Inter Galactic Medium contains:
  - Magnetic fields (regular + random) are highly speculative and range from 2¢10^{-6} nT (20pG) to 10^{-4} nT (1nG).
  - Very little matter (p, He, and a few electrons):
    - 1 proton / m^3
  - Electromagnetic radiations:
    - 413 CMB photons per cm^3
    - Also IR, radio photons…
  - Neutrinos:
    - Mostly C^{o}B neutrinos (decoupled when universe was only 2" old!)
    - Today 1.95 K i.e. 1.7¢10^{-3} eV
    - 336 ^{o} (all species) per cm^3

+ Many mysterious dark matter WIMPs ...



The earth atmosphere

- An evident characteristic of the atmospheric medium is that of being inhomogeneous.
  - Its density, diminishes six orders of magnitude when the altitude above sea level passes from zero to 100km, and another additional six orders for the range 100km to 300km.
  - Although up to ~100km, the composition is nearly constant: 78.47% N, 21.05% O, 0.47% Ar and 0.03% other elements.
  - It follows a quasi exponential profile ("quasi" because T is not quite constant!)

$$\rho(h) = \rho_0 e^{-gMh/RT}$$

The earth atmosphere

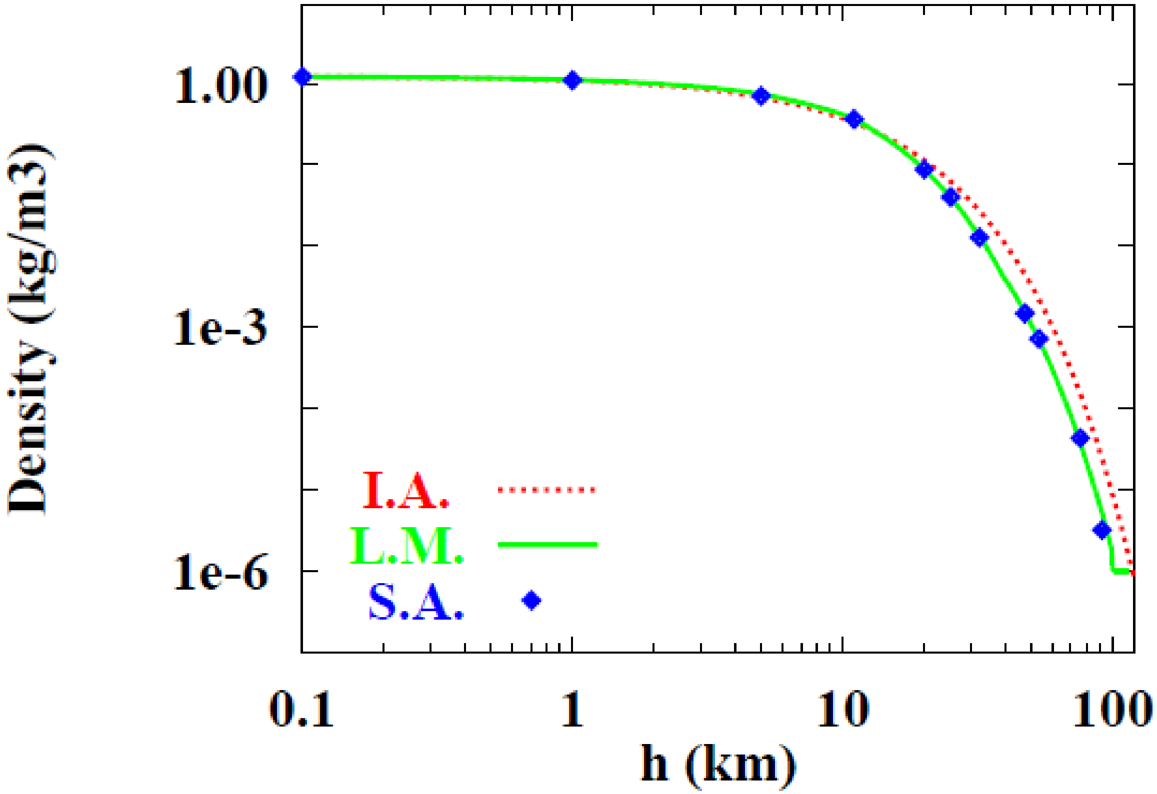


Figure 2.3. *Density of the air as a function of the vertical altitude. The dots represent the US standard atmosphere data [14], while the full green line corresponds to Linsley's model [16] and the dashed red one to the isothermal atmosphere*
$\rho(h) = \rho_0\, e^{-gMh/RT}$ *with* $\rho_0 = 1.225\ \mathrm{kg/m^3}$, $M = 28.966$ *and* $T = 288\ K$.

$$\rho(h) = \rho_0 e^{-gMh/RT}$$

$$\text{Scale height } gM/RT \approx 9km$$

The earth atmosphere

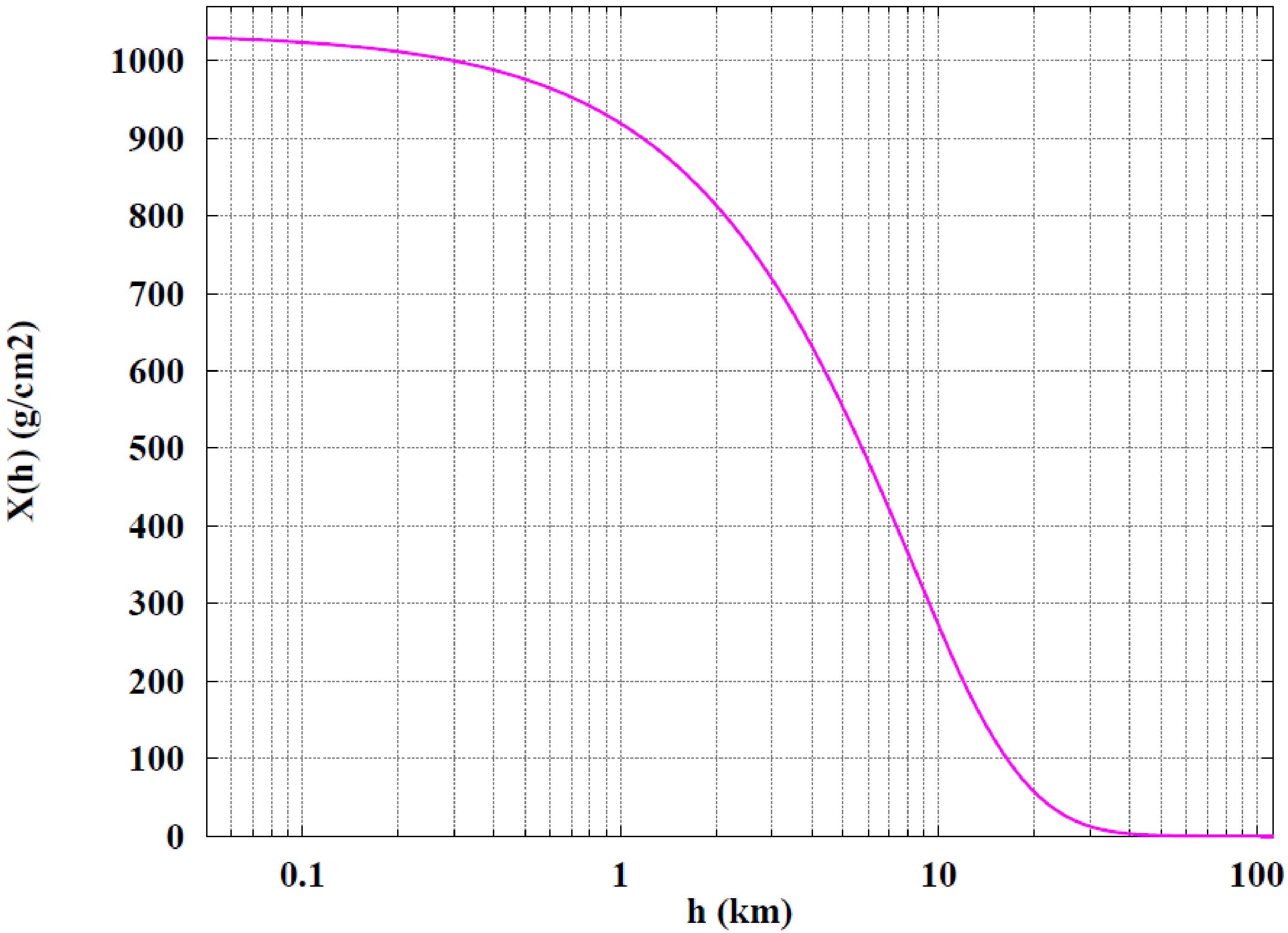
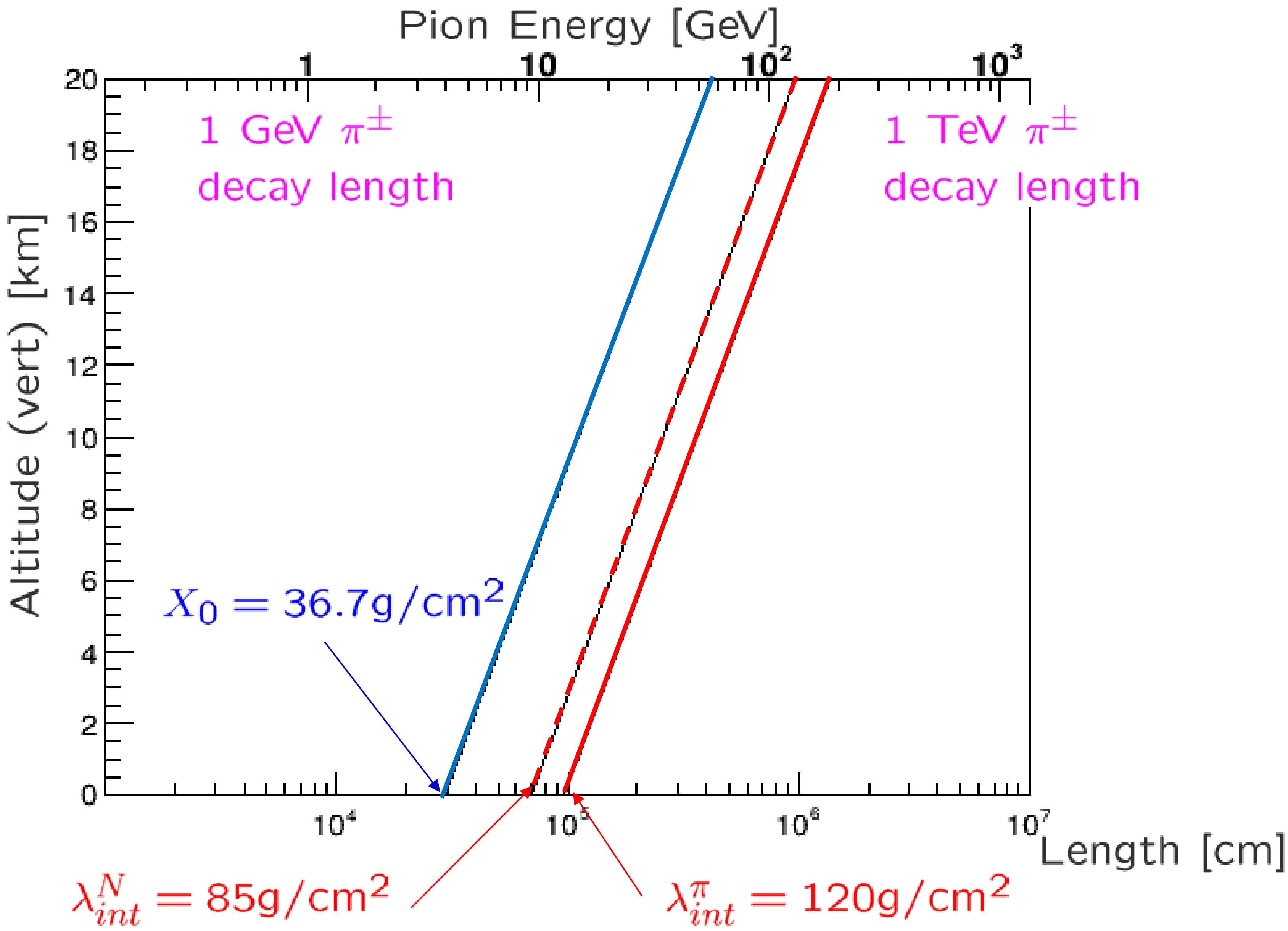


Figure 2.4. *Vertical atmospheric depth, X_v, versus vertical altitude over sea level, h, accordingly with Linsley's model [16].*

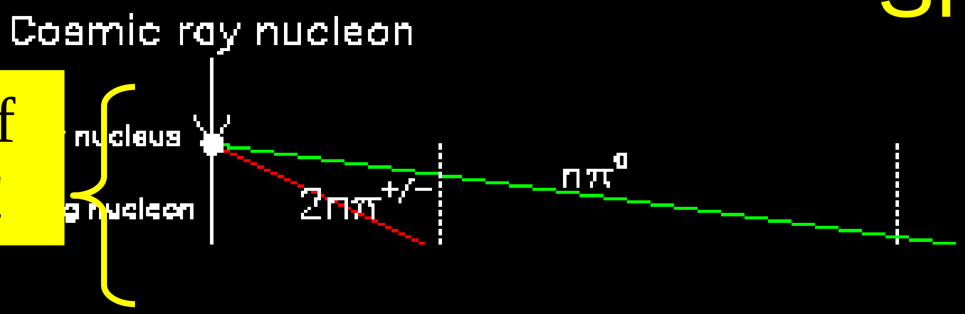
The earth atmosphere

- In terms of particle/radiation interaction with matter, the atmosphere is:
  - A total of ¼ 1000 g/cm^2 at see level
  - So 1 atm ~ 12 interaction lengths ($¸_N$ ¼ 85 g/cm^2)
  - A vertical proton first interacts at h ~ 15 km
  - One radiation length (at 1 atm) X_0 = 36.6 g/cm^2 ¼ 300 m
  - One Moliere radius (at 1 atm) is $½_M$ ¼ 78 m
  - The Lorentz factor for a muon produced at h = 10 km to reach ground before decaying is ¡ > 15 (> 1.6 GeV)
  - Critical energy (EM) E_c = 84,2 MeV

Interaction and radiation lengths in atmosphere



Shower development



Physics well out of reach of colliders !

At each step energy is shared by more numerous particles

"Hadronique" shower

"Electro-Magnétique" shower

$\lambda_{interaction} < \beta\gamma c\tau_{decay}$

$\lambda_{radiation} < \lambda_{ionisation}$

Maximum of développement

Critical Energy E_c

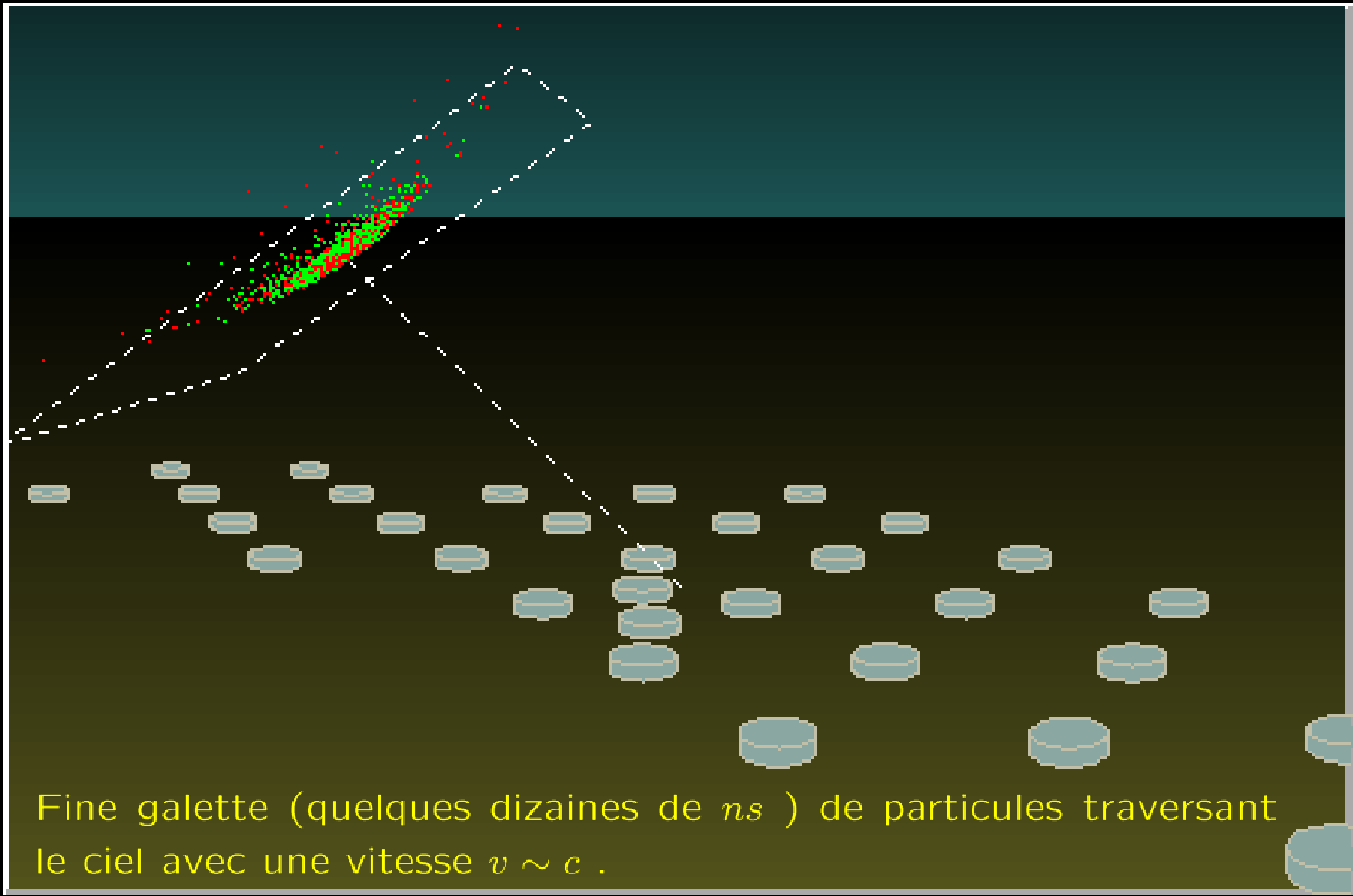
No more multiplication, decrease by decay and energy loss

$\lambda_{interaction} > \beta\gamma c\tau_{decay}$

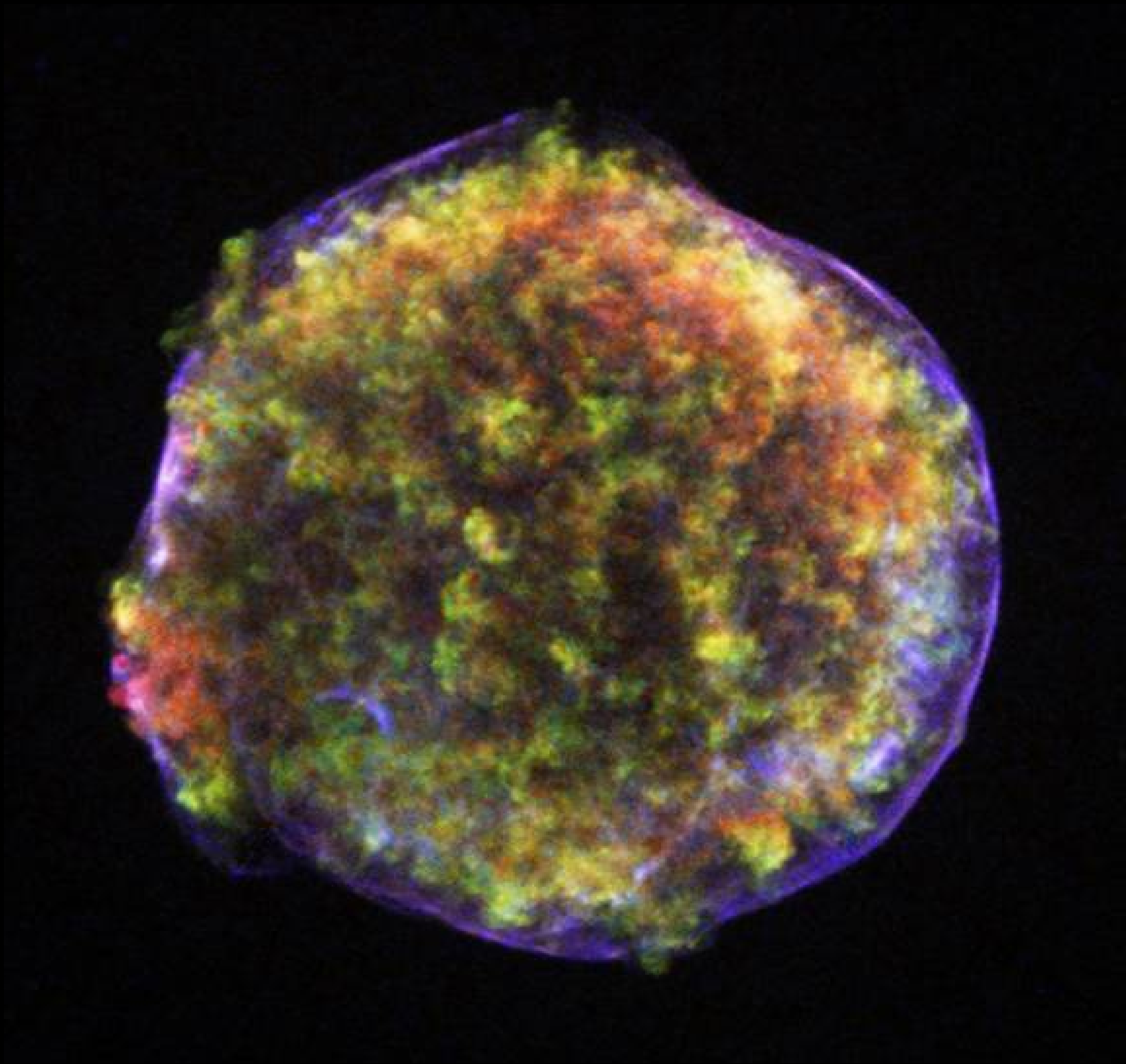
$\lambda_{radiation} > \lambda_{ionisation}$

At ground, essentially m§ ge§

Structure en temps



Astrophysical Sources
Cosmic Accelerators



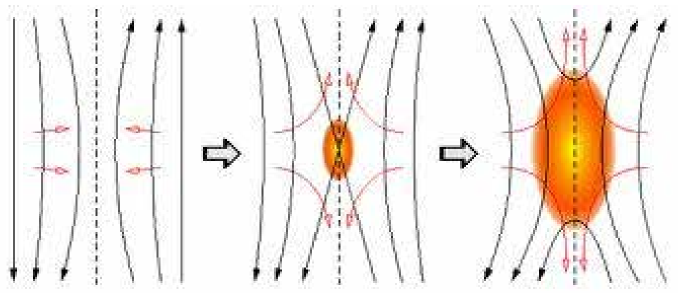
General ideas on acceleration

- Take the necessary energy somewhere...
  - Kinetic energy:
    - Translation: shock waves, moving clouds...
      ® Fermi acceleration
    - Rotation : pulsars, black holes, neutron stars
  - Gravitational energy
    - via accretion (® jets...): accretion disks (divers)
  - Electromagnetic (EM)
    - From turbulence, from compression, or from rotating magnets...
- In fine, charged particles interact
  with EM fields: f = q (E + v´ B)
- Remember:
  Astrophysical shocks are collisionless.

® Energy transfert through EM fields !

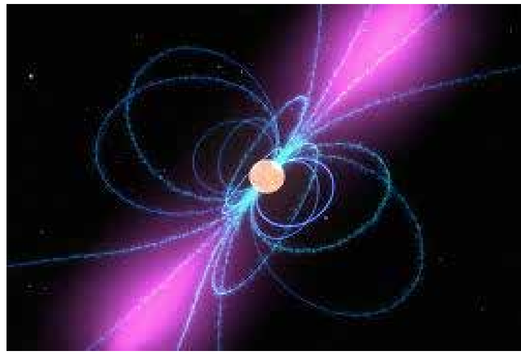
E and B fields in Universe

* In ISM as on earth, $< E > \approx 0$
  - ISM is neutral or conducting



* Transient electic fields:
  - Magnetic re-connections (e.g. solar flares...)
  - EM waves

* Producing E fields in EM "engines"
  - Astrophysical dynamos, induction machines



* Magnetic fields :
  - $\varepsilon_B \approx 1eV/cm^3 \approx \varepsilon_{optique} \approx \varepsilon_{CMB} \approx \varepsilon_{CR}$!!!
  - Astrophysical plasmas :
    ISM, stars, accretion disks, IGM, jets, etc...

Magnetic field production

- Large scale movements of ionized media
  - ® generating magnetic fields, magnetized clouds...
- Turbulence in interstellar medium
  - ® Magnetic turbulence, inhomogeneous B fields, plasma waves...
- Hydro and MHD instabilities
  - – e.g. Rayleigh-Taylor in supernova remnants
- "Streaming" instabilities
  - – CR generates waves in a magneto-actif plasma
  - ® creating the conditions for their own diffusion

Magnetic field production

In many cases, **equipartition** can be reached

- for ex: behind a shock wave :
  thermal en. ~ kinetic en. ~ magnetic en.

Þ Energy exchange between macroscopic structures and individual particles

Þ individual particles may reach very high energies!

Magnetic fields and acceleration !

- **How is it possible at all?**
  magnetic fields don't work ! ($\vec{F} \perp \vec{B}$)

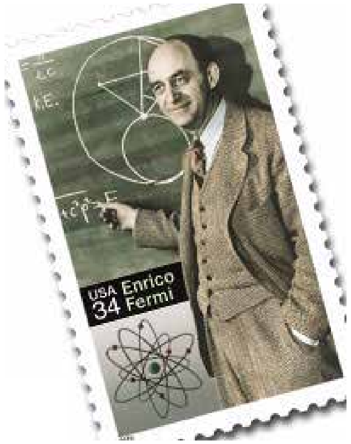
- Well, variable $\vec{B}(t)$ fields do ! (example: *Betatron*)
  $$\vec{\nabla} \times \vec{E} = -\frac{\partial \vec{B}}{\partial t}$$

- In a different reference frame,
  a pur $\vec{B}$ field is feeled as a $\vec{E}$ field...
  $\vec{E}' = \vec{v} \times \vec{B}$ (for $v \ll c$)

- In principle, one can always identified an effective $\vec{E}$ field that works, but the description in terms of $\vec{B}$ fields is often simpler (and more physical !)

® Acceleration by "change of reference frame"

Principle of Fermi acceleration



The Ingredients :

- **A magnetic field** $\vec{B}$
  $=$ with a regular component $\vec{B}_0$
  $+$ and irregular component $\delta\vec{B}$
- **A plasma** i.e. a good electrical conductor :
  $\vec{E} + \vec{u} \times \vec{B} = 0$ and $|E| \approx 0$
  $\Rightarrow$ the magnetic field is "frozen" and moves with the plasma (Alfven).
- **A CR population** coupled to the medium via the magnetic field $\vec{B}$. They scatter on the field irregularities. This diffusion processes are **collisionless** i.e. they conserve the particle energy. The MHD or Alfven waves act as massive scattering centers (recoilless).

Fermi 1949 :

- first hypothesis of converging movements of MHD perturbations
  $\Rightarrow$ "first order" acceleration, but where ?
- second more realistic hypothesis at that time: random mouvement of interstellar gas clouds (observed) or MHD perturpations
  $\Rightarrow$ "second order acceleration.

Where to accelerate

- At creation :
  - For example: e^- extracted from the surface of a neutron star by an intense E field.
- Within the source neighborhood :
  - For example: Fermi acceleration in plasma shocks in a SNR.
- During transport:
  - "reacceleration" by shock waves and excitation of Alfven waves during diffusive transport in the Galaxy.

Power laws and stochastic processes

- The power laws observed in differential energy spectra follow naturally from cyclic acceleration mechanisms with constant energy gain and constant escape probabilities:

  - Initial energy: E_0
  - Energy gain at each cycle: $\Delta E = \varepsilon E$
  - Particle energy after n iterations: $E_n = E_0(1+\varepsilon)^n$
  - Escape probability from the acceleration zone: P_{esc}
  - Probability to remain in the acceleration zone: $(1 - P_{esc})^n$

Power laws and stochastic processes

- Particle energy after n iterations: $E_n = E_0(1+\varepsilon)^n$
- Probability to remain in the acceleration zone: $(1-P_{esc})^n$

Number of iterations to reach an energy E:

$$n = \frac{\ln(E/E_0)}{\ln(1+\varepsilon)}$$

Proportion of particles accelerated up to an energy equal or greater than E :

$$N(\geq E) = N_0 \sum_{m=n}^{\infty} (1-P_{esc})^m = N_0 \frac{(1-P_{esc})^n}{P_{esc}}$$

thus :

$$\frac{\ln(P_{esc}N/N_0)}{\ln(1-P_{esc})} = n = \frac{\ln(E/E_0)}{\ln(1+\varepsilon)}$$

eliminating n:

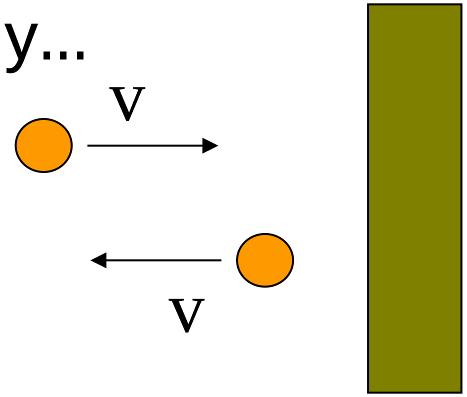
$$N(\geq E) \propto \left(\frac{E}{E_0}\right)^{-\gamma}$$

with $\gamma = \frac{-\ln(1-P_{esc})}{\ln(1+\varepsilon)} \approx \frac{P_{esc}}{\varepsilon} = \frac{1}{\varepsilon}\frac{T_{cycle}}{T_{esc}}$

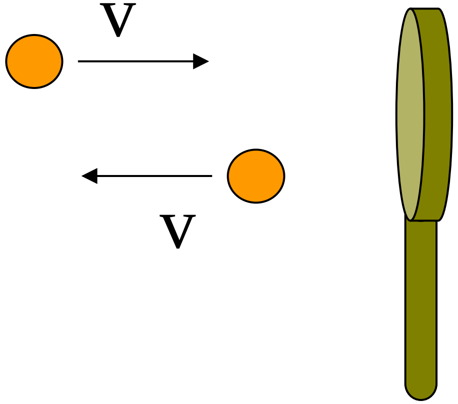
→ Power laws are natural !

A small analogy…

- A tennis ball bouncing on a wall
  - neither gain nor loss of energy…



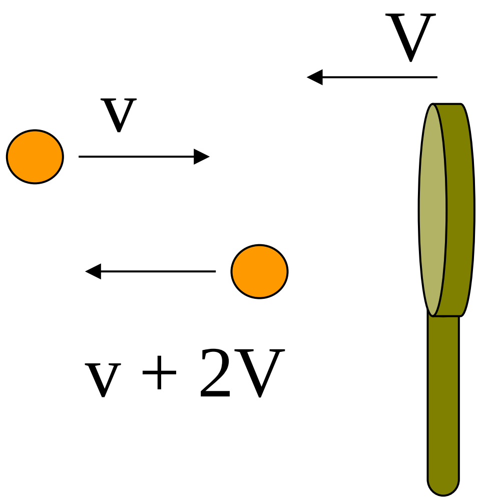
bounce = speed unchanged



Same thing with
a motionless racquet…

Then how does one accelerate a tennis ball ?!

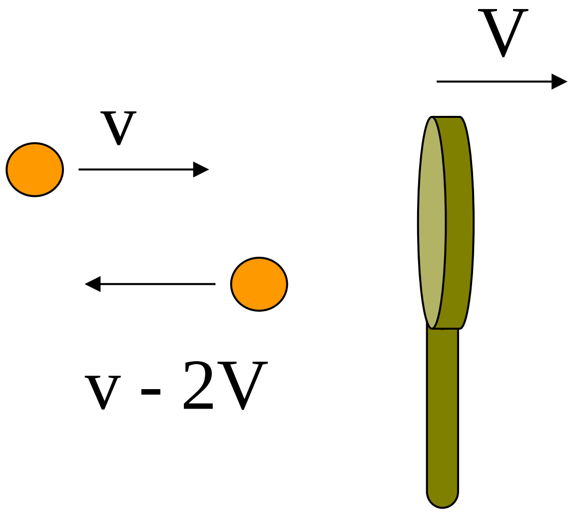
- Moving racquet
  - Neither gain nor loss of energy… in the racquet reference frame !



Speed unchanged with respect to the racquet

® acceleration through a change of reference frame

- A drop shot:

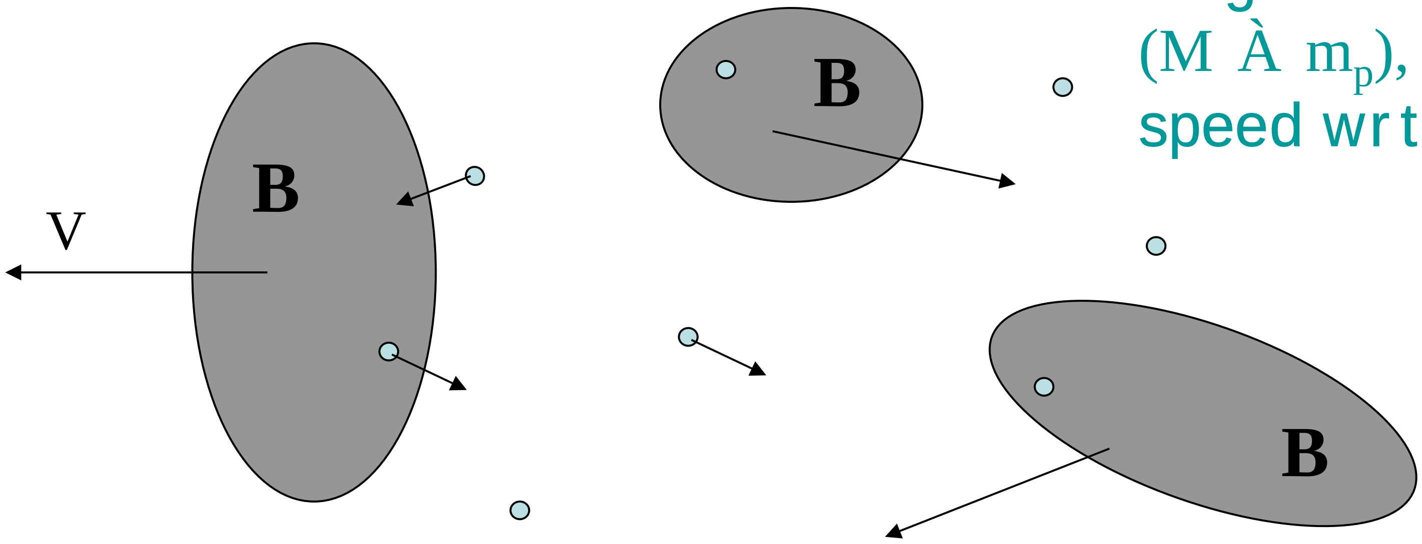


Particle deceleration !

Fermi Acceleration

- Ball ® charged particle
- Racquet ® "magnetic mirrors"

Magnetized clouds (M À m_p), with a typical speed wrt ISM ~10km/s



- Magnetic inhomogeneities or plasma waves also work...

The essence of stochastic de Fermi acceleration

1 When a particle bounces on an **incoming** magnetic mirror, in a **head-on** collision, it **gains** energy.

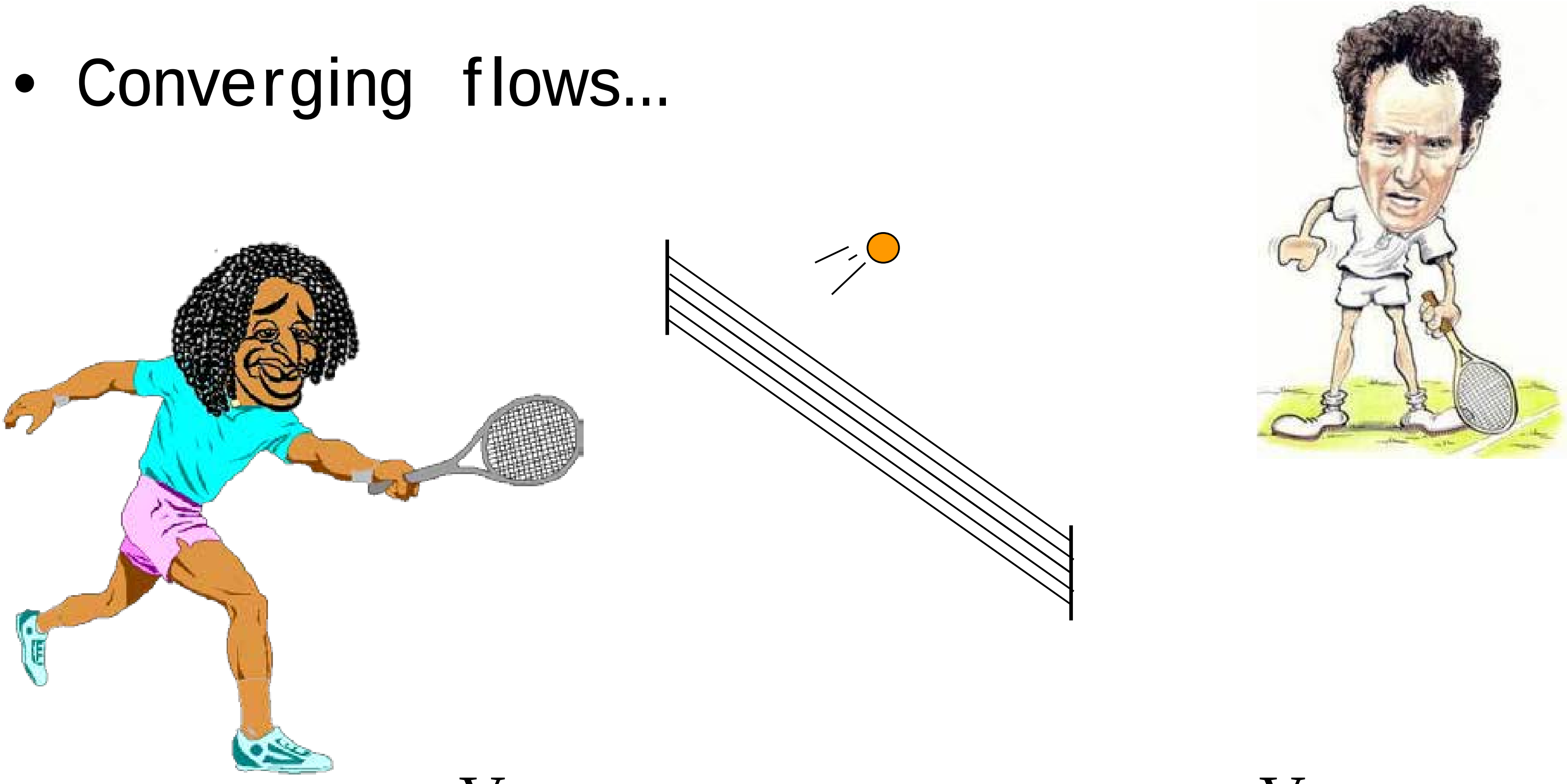
2 When a particle bounces on a **receding** magnetic mirror that it catches back, it **loses** energy.

3 Head-on collisions are **more frequent** than receding collisions.

Þ Net energy gain in average (stochastic process)

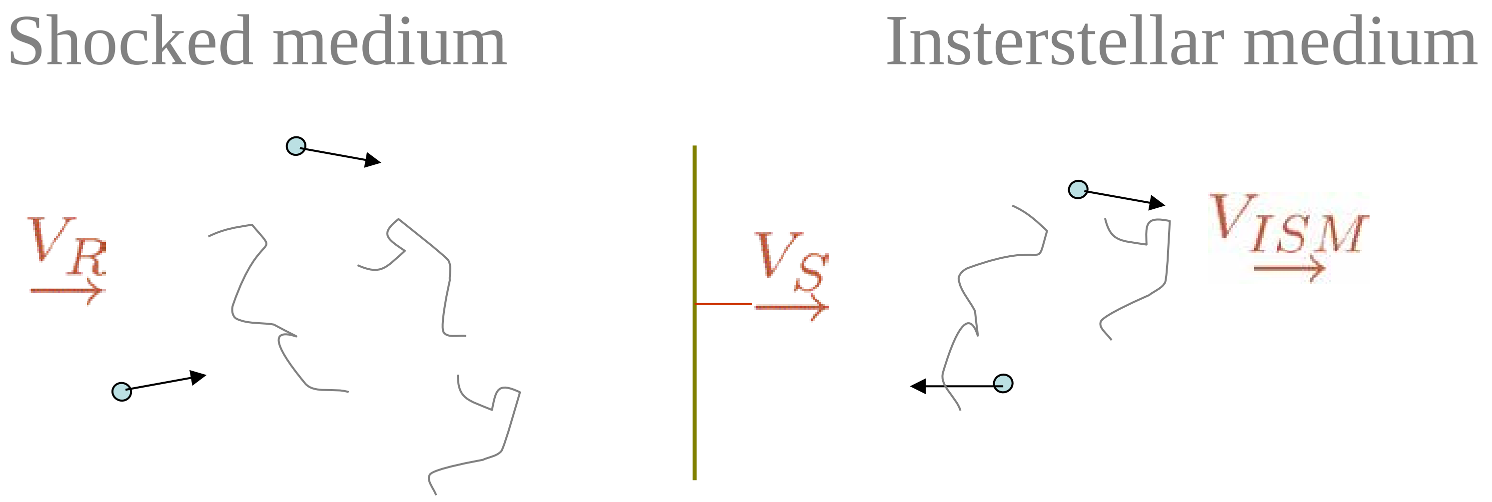
Add a second player...

- Converging flows...



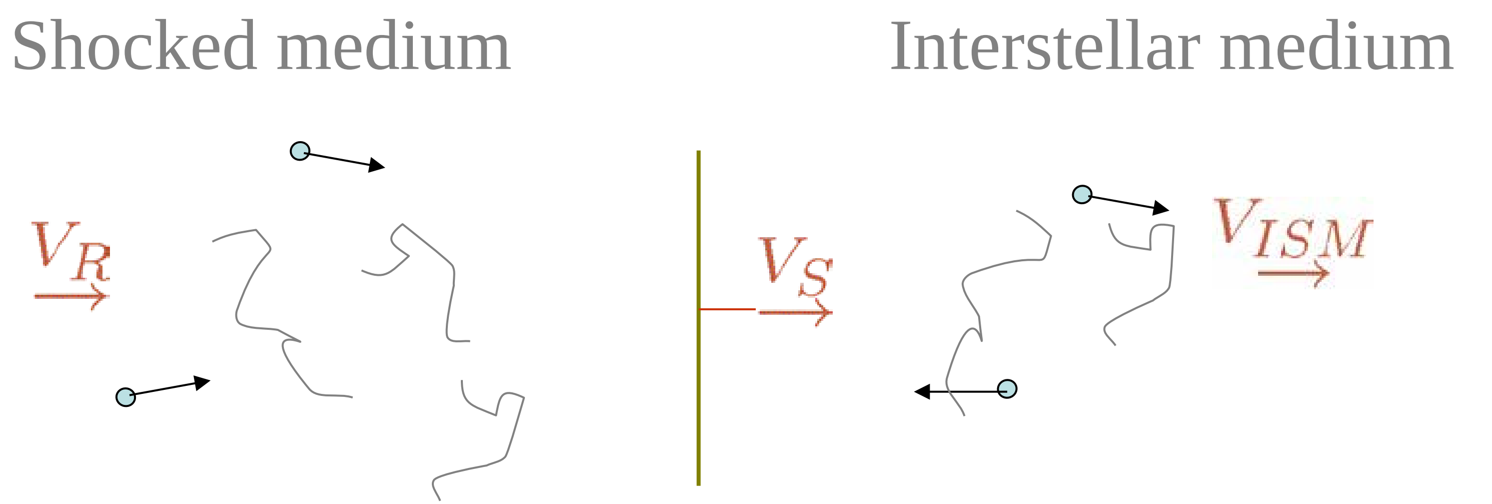
Shocks hydrodynamics

- Shock waves (e.g. supernova explosion) : expending plasma flow with a speed V_R much larger than the sound speed in the interstellar medium (ISM).



Shocks hydrodynamics

- Shock wave:

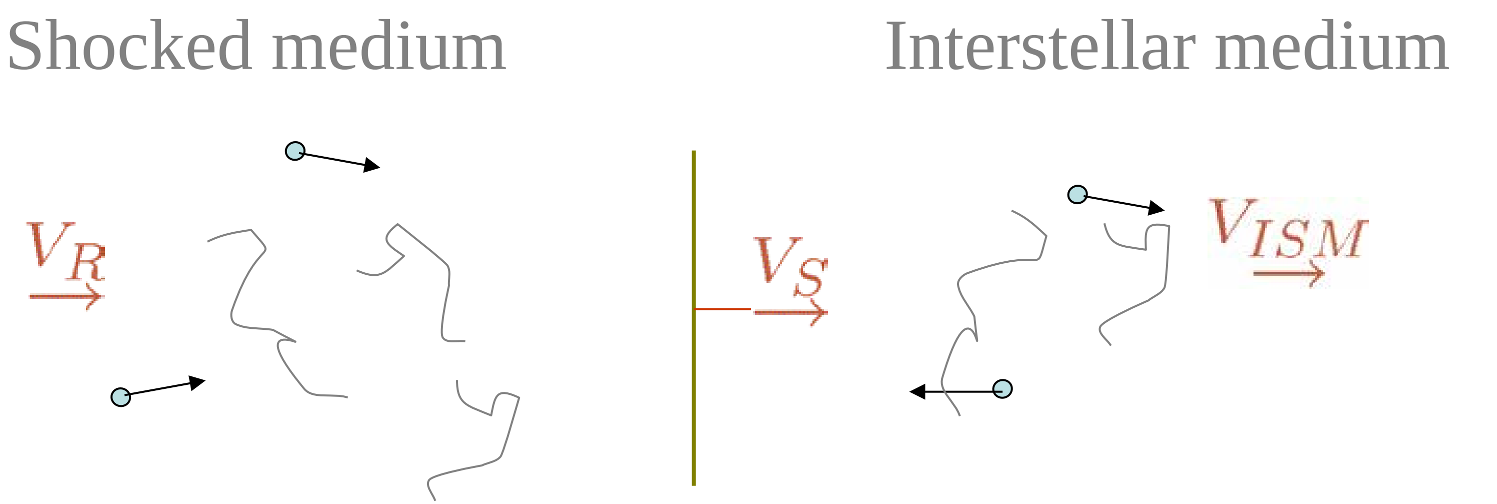


- The shock moves at a speed V_S which depends on V_R and the specific heat of both media.

- For an ionized ISM: $$V_S \approx \frac{4}{3} V_R$$

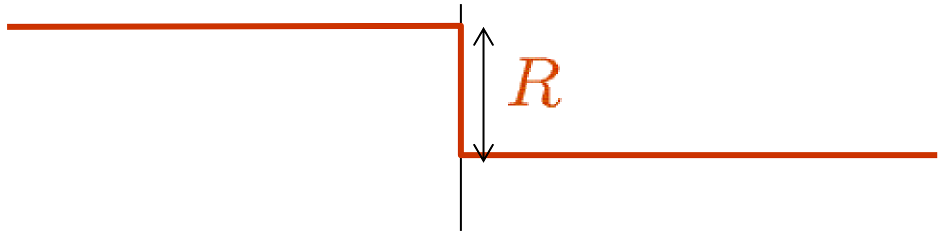
Shocks hydrodynamics

- Onde de choc :



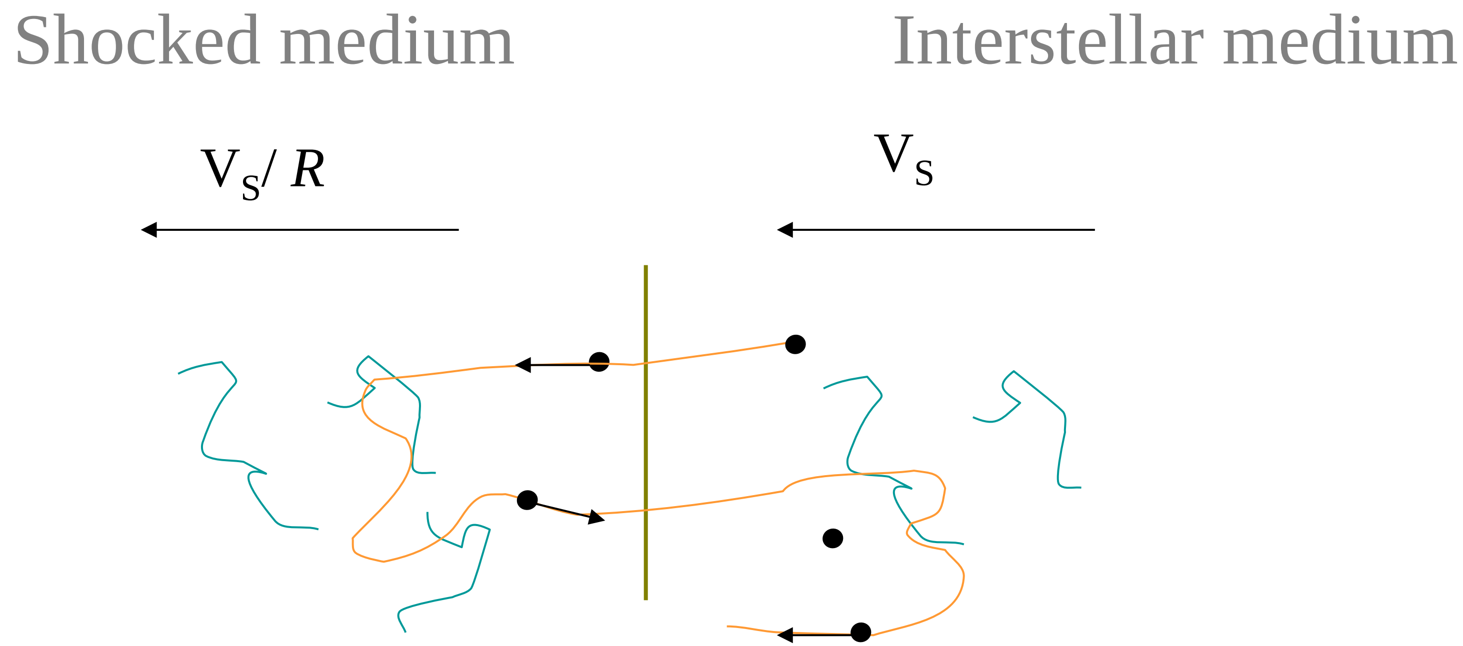
- The shock intensity is characterized by the compression factor:

$$R = \frac{V_S/V_R}{V_S/V_R - 1} \approx 4$$



Shocks hydrodynamics

In the shock frame

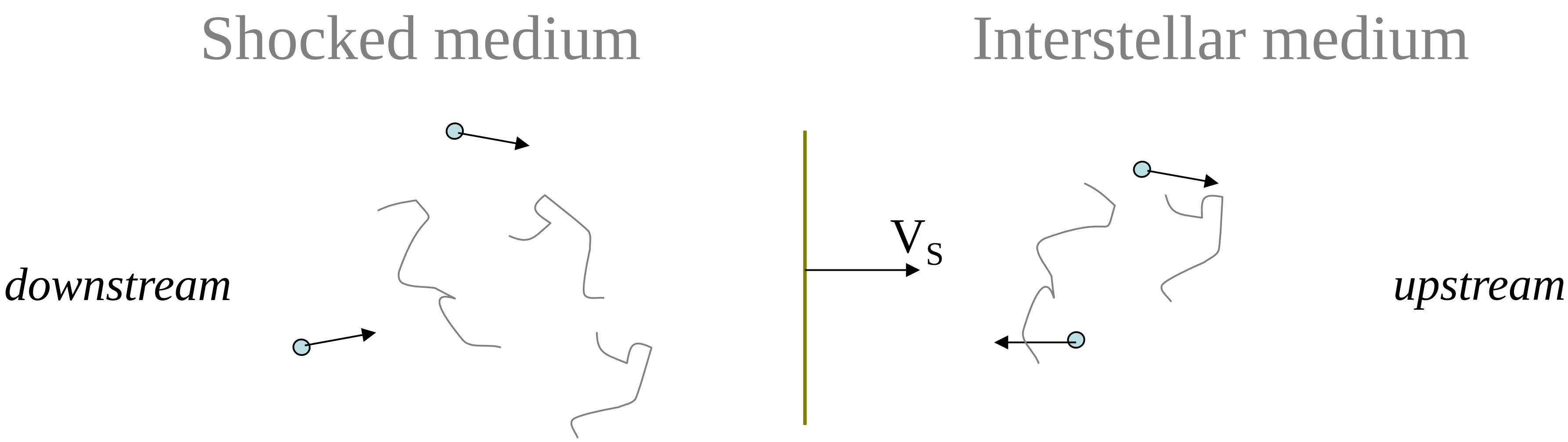


- In the shock frame, the upstream (non-shocked) medium flows toward the shock at a speed V_S and the downstream (shocked) medium flows away with a speed reduced by the compression factor (mass flow conservation) :

$$V_S / V_d = R \approx 4$$

Acceleration diffusive par onde de choc

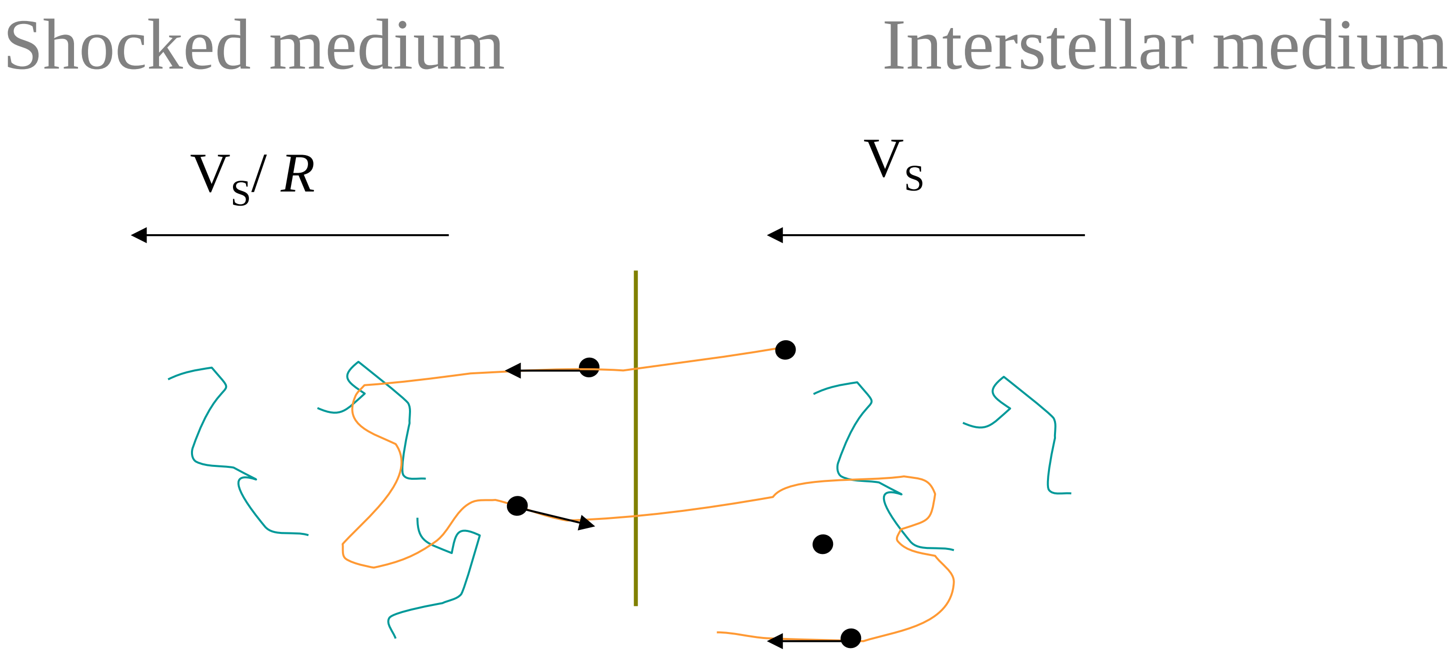
- Onde de choc (e.g. explosion de supernova)



- Magnetic wave generation:
  - **Downstream** : by the shock (compression, turbulence, hydro and MHD instabilities, shear, etc.)
  - **Upstream** : by the accelareted cosmic rays themselves !

® 'isotropization' of the distribution (in the local frame)

A win win process !



In the shock frame

- At each shock crossing, one way or the other, the particle hits a "magnetic wall" with a relative speed:

$$V = (1 - 1/R) V_S$$

® only head-on collisions...

Summary on acceleration

- Acceleration from interaction with fields
  - **E** field: e.g. induced by spinning magnets such as neutron stars (pulsars) or black holes...
  - **B** field: inhomogeneous moving fields
  - MHD waves
- Acceleration by reference frame transformation
  - Fermi stochastic acceleration (2nd order)
  - Diffusive shock acceleration diffusive (1st order)
- Power law are natural
  - Fermi type process **($\Delta E \propto E, P_{ech}$)**
  - Universal power law for non relativistic shocks **($N(E) \propto E^{-2}$)**
- Cosmic rays up to the knee
  - CR power = power of SNe, $\mathbf{E_{max} \approx 10^{14}}$ eV hardly $\mathbf{10^{15}}$ eV

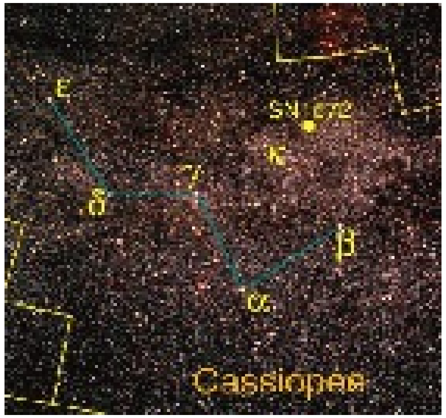
The CR standard model

- Analytic calculations, simulations and observations show that diffusive shock acceleration works !

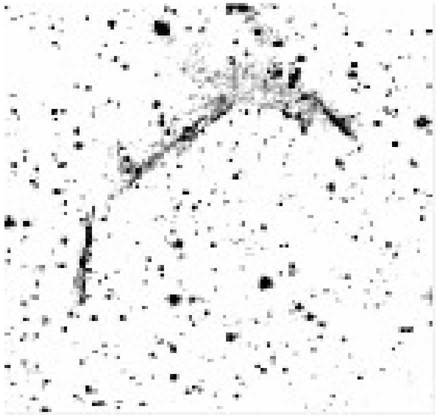
- Supernovae and GCRs
  - Estimated efficiency of shock acceleration :... $10-50\%$
  - Power required to sustain CR energy density: ..... $\varepsilon_{CR} \times V_{conf}/\tau_{conf}$ $\sim 10^{41}\text{erg/s}$!
  - Power injected by SN power in the Galaxy: ....10^{42}erg/s !

® Enough power for Galactic CR

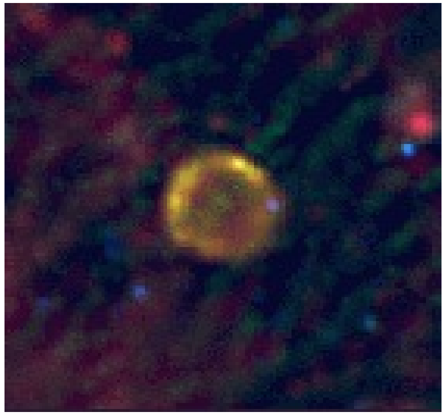
Tycho, 11 November 1572...



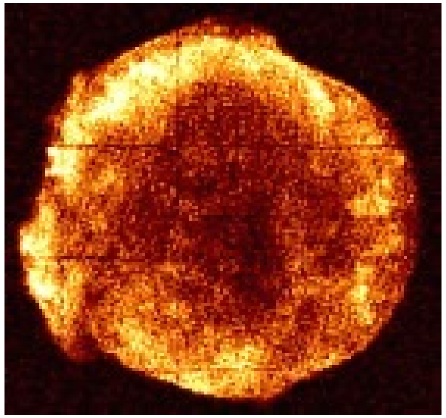
position



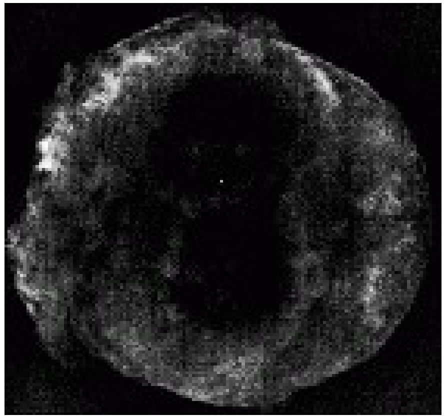
visible



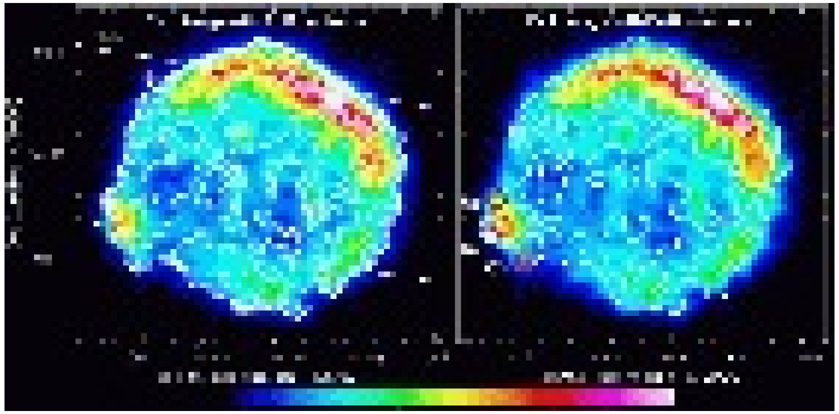
IRAS



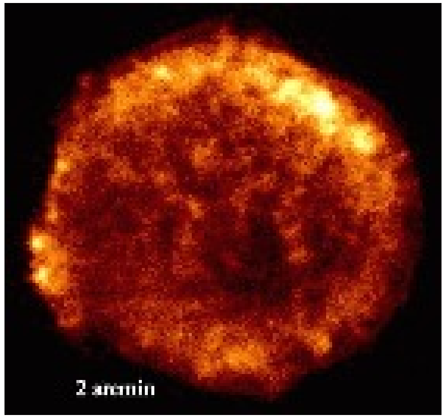
Km (VLA)



6 cm (VLA)



Si K (XMM) Fe K



X (ROSAT)

The CR standard model

- Proposed acceleration site, isolated SNR
  - Supernovæ : ejection of many solar masses of nuclear matter at supersonic speeds (~ 10 000 km/s) following massive star explosion.
  - Formation of a quasi spherical expending shock wave that wipes out the interstellar medium (ionized beforehand by the progenitor's radiation).
  - Total kinetic energy injected by the explosion: 10^{51} erg (= 10^{44} J).
  - Roughly 3 SNe per century witin our Galaxy, which corresponds to an averaged power of 10^{42} erg/s (10^{35}W)
  - SNR are observed at all wavelength.
  - SNe explosion is essential to the Galaxy chemical content: heavy elements enrichment.

The CR standard model

- Shock waves in isolated SNe (SNR)
  - Source composition source ~ interstellar medium + modifications (ionizability, volatility, Z/A effects, ^{22}Ne...)
  - Source spectrum: E^{-2} power law
  - Maximal energy reached: E_{max} ~ 10^{14} eV
- Energetics :
  - Measured flux / speed = CR density
  - CR density ´ mean energy = energy density
  - Energy density ´ confinement volume = total energy
  - Total energy / confinement time = necessary injected power
  - P_{CR} ~ 1.5´ 10^{41} erg/s
- Required efficiency ~ 10-30%...

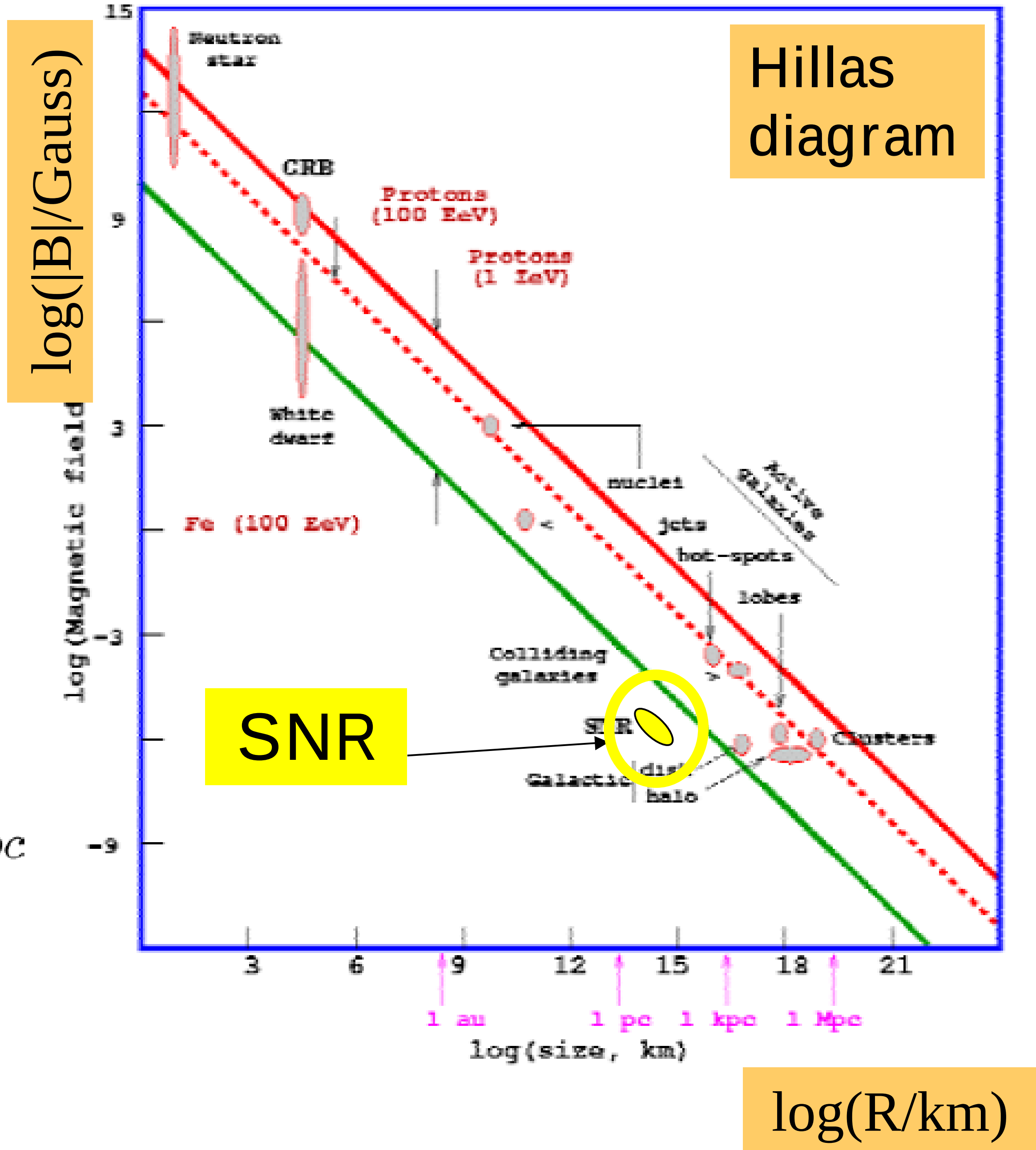
Finite size

Finite size of confinement magnetic field

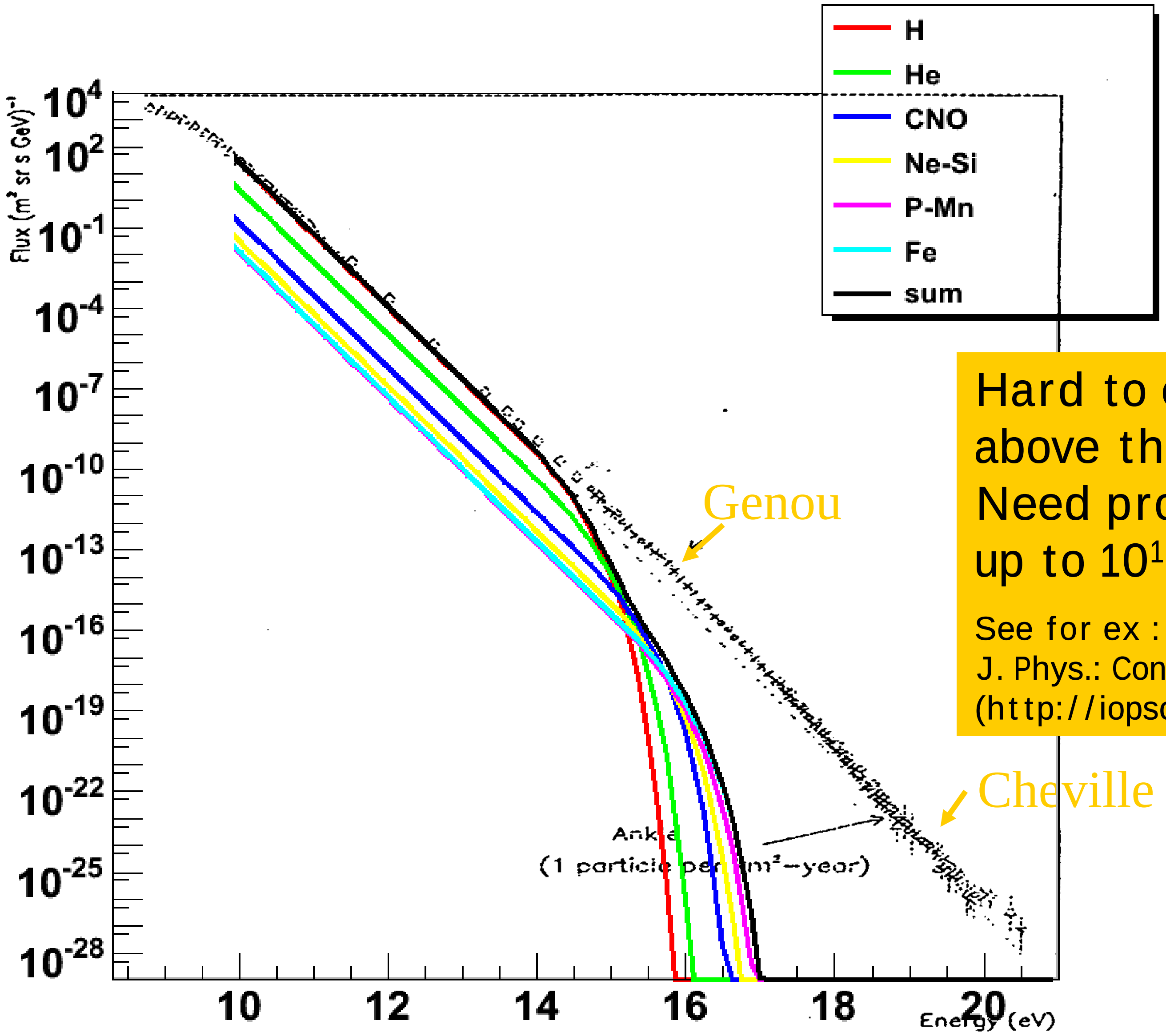
⇕

Larmor radius

$$r_g \leq R$$
$$\Leftrightarrow \quad \frac{p}{ZeB} \leq R$$
$$\Leftrightarrow \quad E \leq Ze \times B \times R$$
$$\Leftrightarrow \quad E \leq (10^{17} eV) Z B_{\mu G} R_{pc}$$



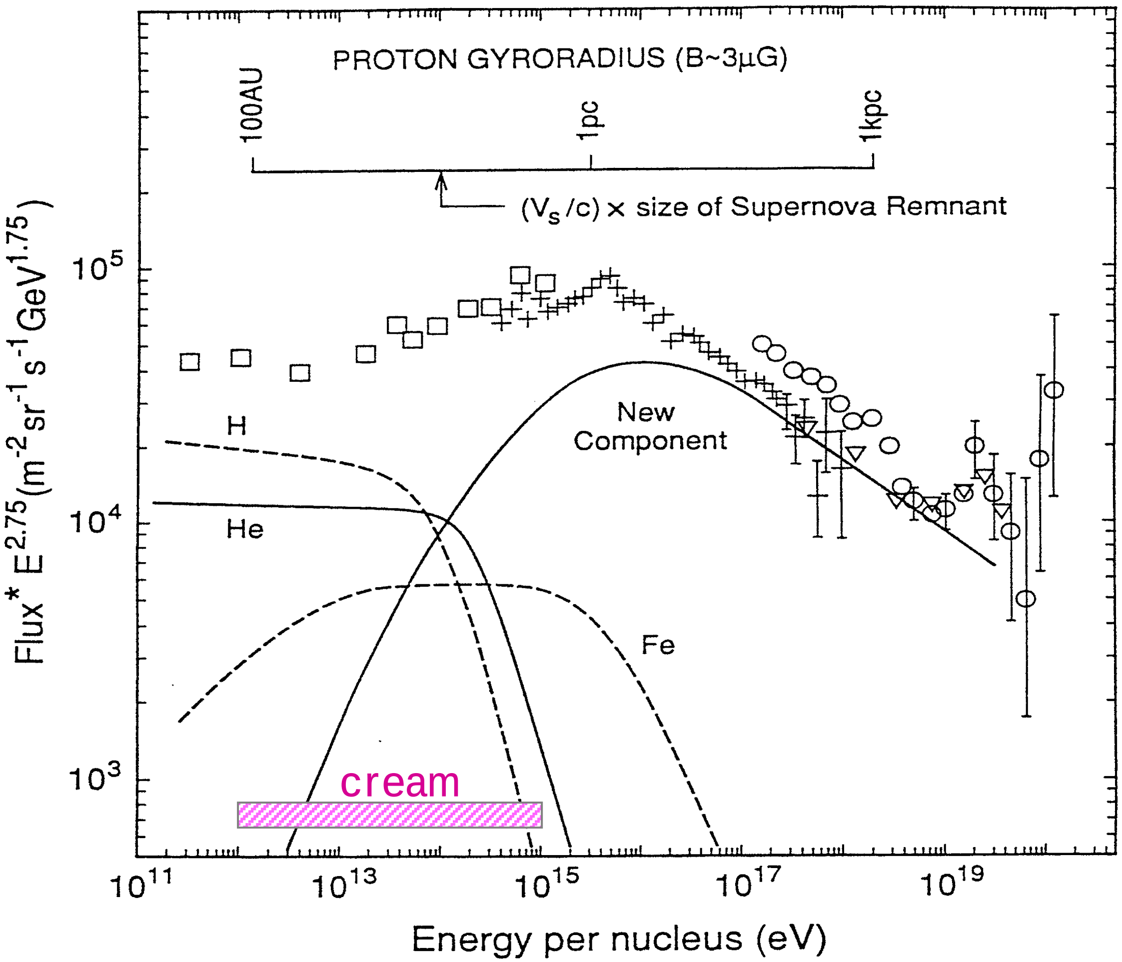
CR SM with Emax / Z



Hard to explain a single power law above the knee.
Need probably another component up to 10^{17}eV

See for ex : M. Hillas
J. Phys.: Conf. Ser. 47 (2006) 168
(http://iopscience.iop.org/1742 6596/47/1/021)

The knee



- Is the knee simply the consequence of an energy cut-off of accelerators (SNR) or is it due to :
  - a propagation effect ?
  - different CR sources ?
  - a physics threshold at $\sim E_{LHC}$?

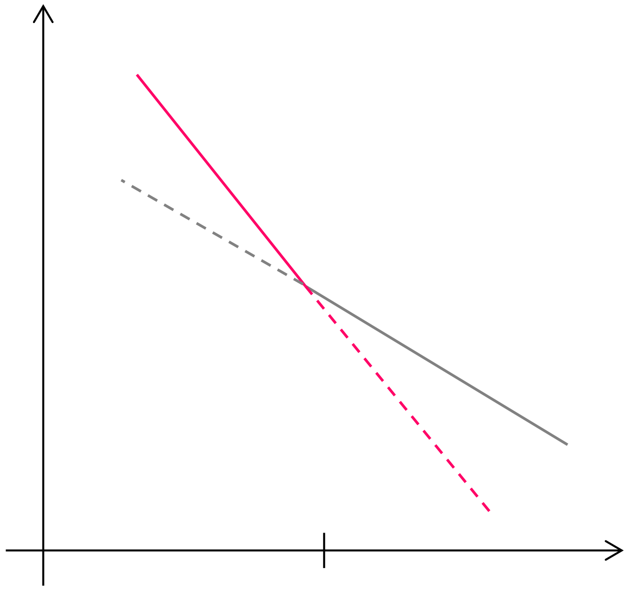
_ The CR SM implies a change in composition at the knee energy.

SNR energy limit : $E_{max} \sim Z \bullet 10^{15}$ eV

Data:
- Proton4 sat
- Akeno
- Yakutsk
- Haverah Pk

Explaining the ankle is much easier...

- Two components with two different slopes...

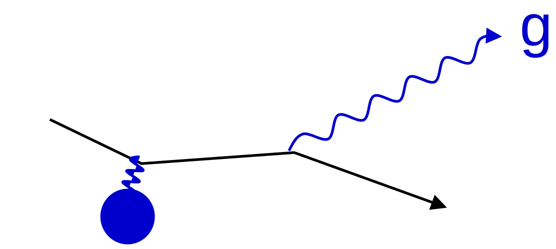


- For exemple, galactic and extragalactic...

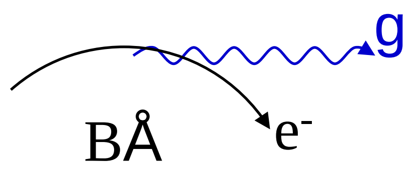
DIFFUSE GAMMA-RAY SOURCE

VHE gamma-rays sources

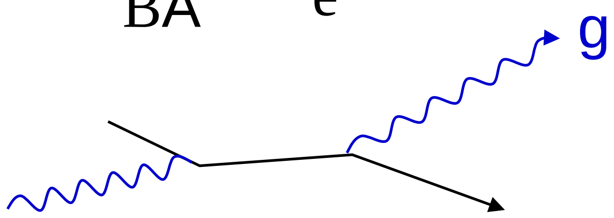
- °-rays production processes:
  - Electro-Magnetic processes :
    - Bremsstrahlung



    - Synchrotron radiation



    - Inverse Compton



  - Hadroniques processes:

$$
\begin{aligned}
(p,N) + (N\gamma) \rightarrow \;\; & \pi_0 + \dots \\
& \hookrightarrow \gamma\gamma
\end{aligned}
$$

Hadronic production of gamma-rays

Source function [TK Gaisser]

$$q_k(E_k, r) = \int \frac{d\sigma_{i\to k}(E_k, E_i)}{dE_k} \left(\frac{c\rho(r)}{m}\right) \left(\frac{4\pi}{c}\phi(E_i)\right) dE_i$$

Fluxes on earth (neutrals)

$$\phi_{k=\gamma} = \frac{dN_k}{dA dE_k d\Omega} = \int \frac{q_k(E_k, r)}{4\pi r^2} d^3 r = \int_0^{r_{max}} \frac{q_k(E_k, r)}{4\pi} dr d\Omega$$

Scale invariance : $E_\gamma / E_i = Z$ (Z indep. of E)
(for $E_k \gg m_i/2 = m_\pi/2 \approx 70 MeV$)
(system mass scale).

$\Rightarrow \phi_k \propto \phi_i$

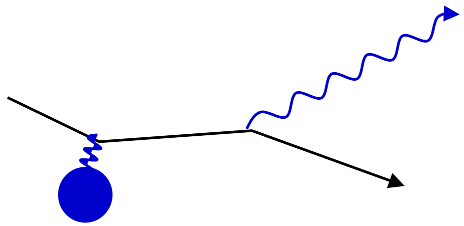
Parent spectral density power law $\phi_i \propto E_i^{-\alpha}$
[Gaisser,Halsen,Berezinsky,Stanev]

$$\phi_\gamma = 4\pi\rho \left[\frac{\sigma_{inel}}{m_N}\right] \left[\frac{2Z_{N\to\pi^0\to\gamma}}{\alpha}\right] \phi_{CR}(E)$$

$$\frac{\phi_\gamma}{\phi_{CR}} \approx 6\times 10^{-4} \times \left(\frac{\rho R}{[g.cm^{-2}]}\right)$$ $\rho R =$ target colonm density

Bremsstrahlung

$$\frac{d\sigma_{e\to\gamma}(E_\gamma, E_e)}{dE_\gamma} = \frac{1}{E_e}\frac{\phi(z)}{N_A X_0}$$



where $\phi(z)$ is a function of $z = E_\gamma/E_e$

For energies $> 70MeV/c^2$, there is a scaling relation between the γ energy and that of the parent electron. (the only mass scale in the problem is the electron mass, much lower than 70 MeV).

$\Rightarrow$ the differential spectrum of the progenitor electrons $(\phi_e(E_e) \propto E^{-\alpha})$ is transmitted to the γ differential spectrum:

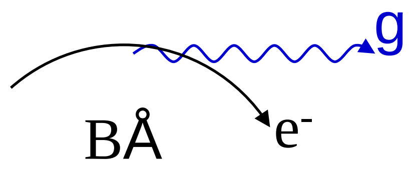
$$\phi_\gamma(E_\gamma) \propto E_\gamma^{-\alpha}$$

For $dN/dE_e = a_e E_e^{-\alpha}$ and $\alpha = 2.7$

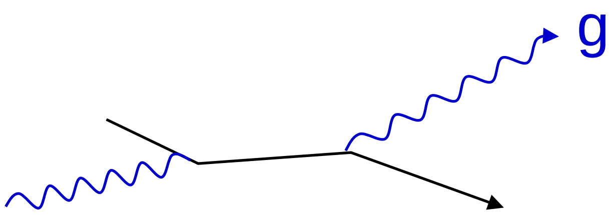
$q_{br} \approx 1.2 \times 10^{-25} a_e n E_\gamma^{-2.7}$ photons GeV^{-1} s^{-1} cm^{-3}

VHE Gamma sources

- Productions of g > 10 MeV
  - Synchrotron radiation negligible.



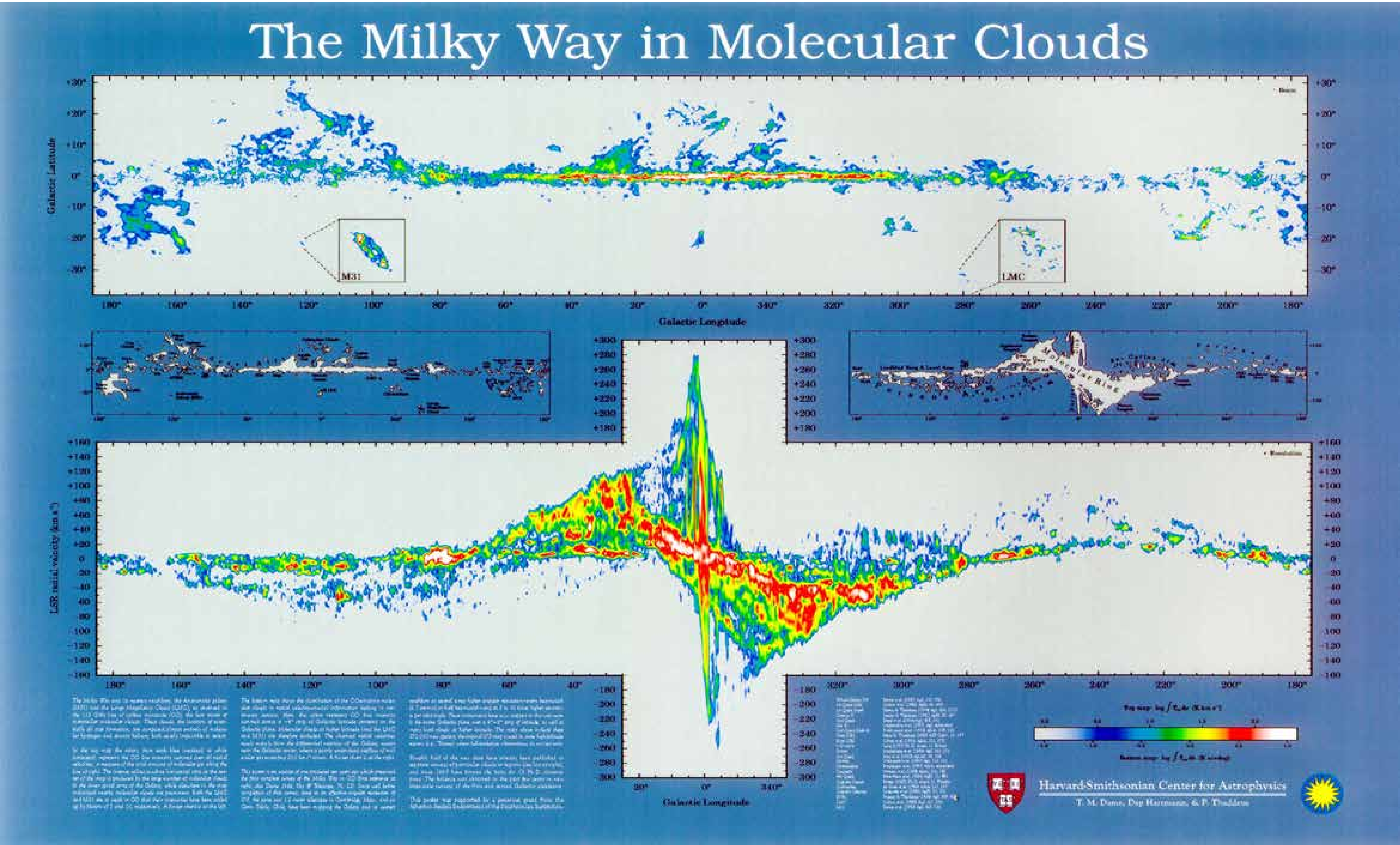
  - Inverse Compton



For $\phi_e \propto E^{-\alpha}$
the IC spectrum is much flatter
$\propto E^{-(\alpha+1)/2}$

but $N_e \ll N_p$ and the electron spectrum drops ↘ above a few GeV
(synchrotron radiotion energy losses)

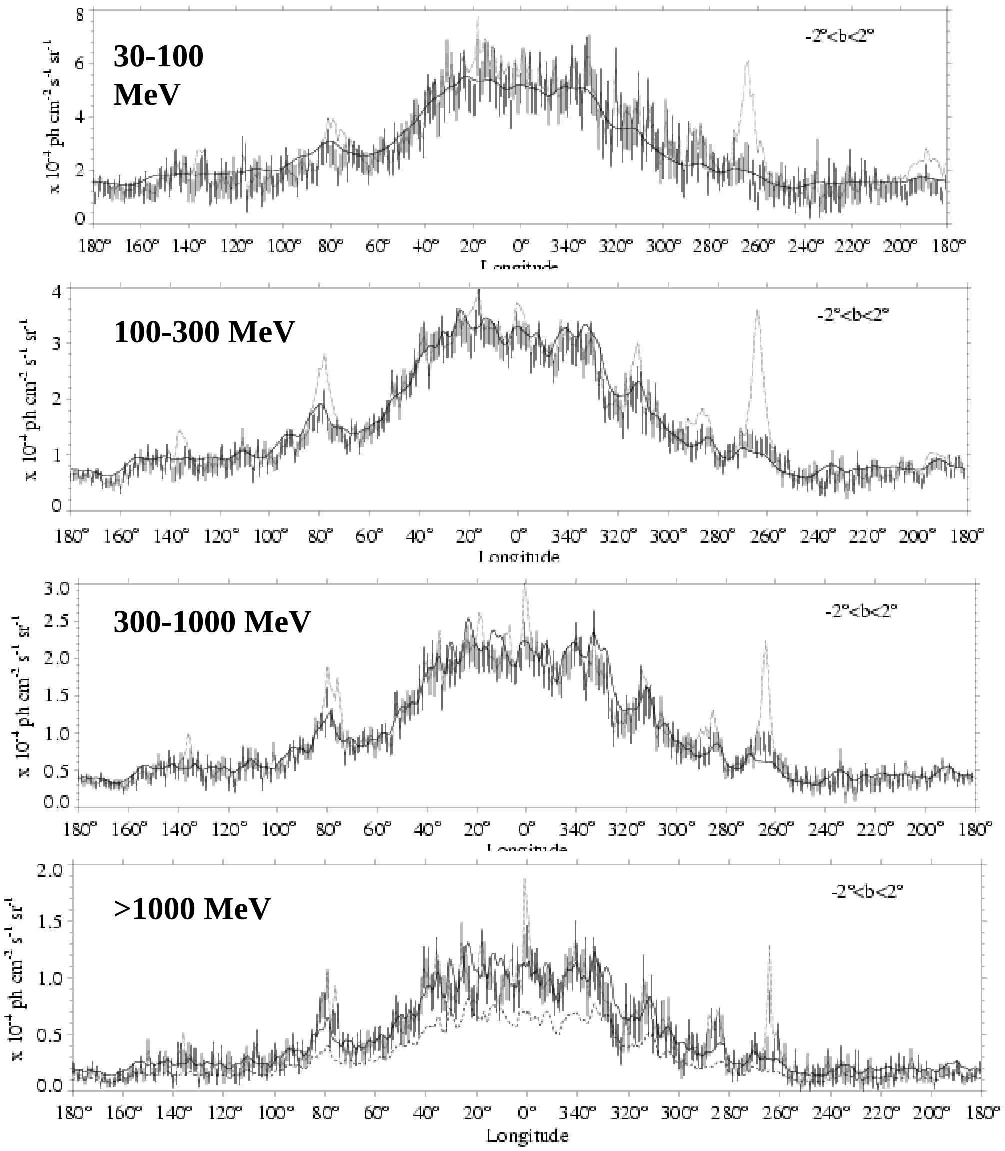
The target…


The Milky Way in Molecular Clouds
Galactic Latitude
Galactic Longitude
M31
LMC
LSR radial velocity (km s⁻¹)
Molecular Ring
Harvard-Smithsonian Center for Astrophysics
T. M. Dame, Dap Hartmann, & P. Thaddeus

Tracking back the CR flux elsewhere in the Galaxy

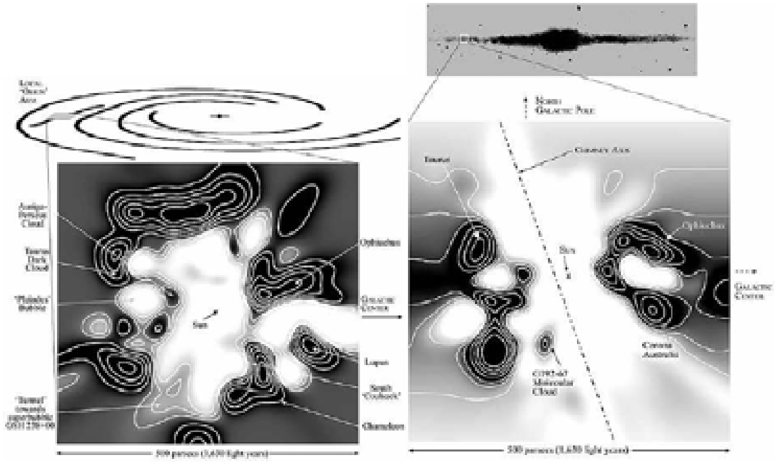
- Knowing the target column density from radio 21cm measurements and fitting the diffuse g flux, one can map the density, spectrum and composition of CR elsewhere in the galaxy.

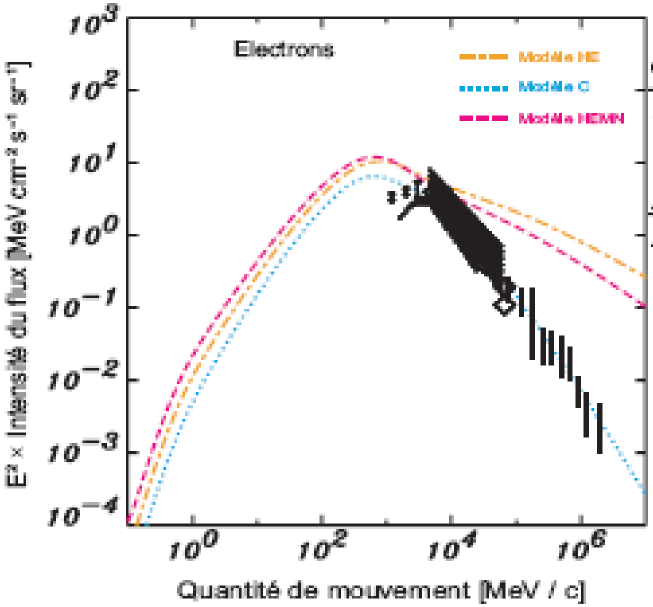
  - S.D.Hunter et al, ApJ 481 (1997) 205
  - M.Pohl and J.A.Esposito ApJ, 507 (1998) 327
  - S.LeBohec et al, astro-ph/0003265
  - Strong, A.W., Moskalenko, I.V., ApJ 509:213-228, 1998

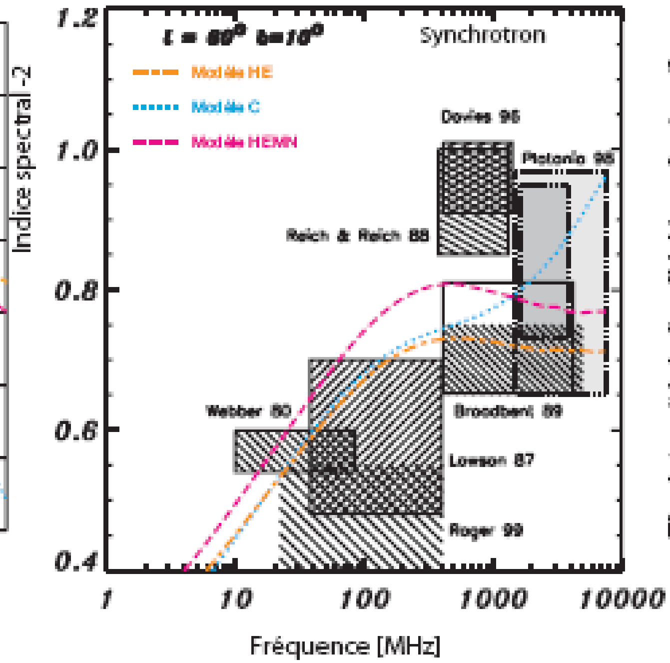


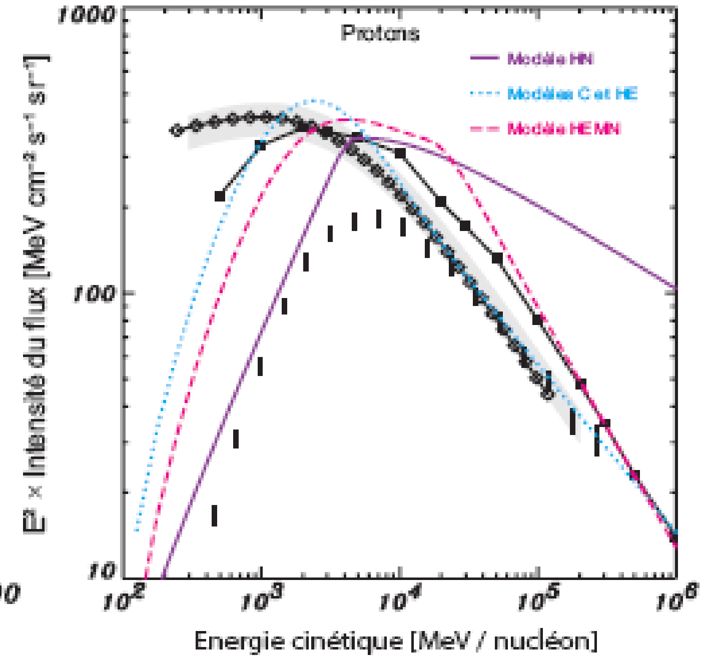
Tracking back the CR flux elsewhere in the Galaxy

- Other constrains
  - Protons local spectrum and flux (altered <10GeV solar magnetic field)
  - Electrons local spectrum and flux (influenced by the local bubble of matter under-density)
  - Radio measurements of the synchrotron emission



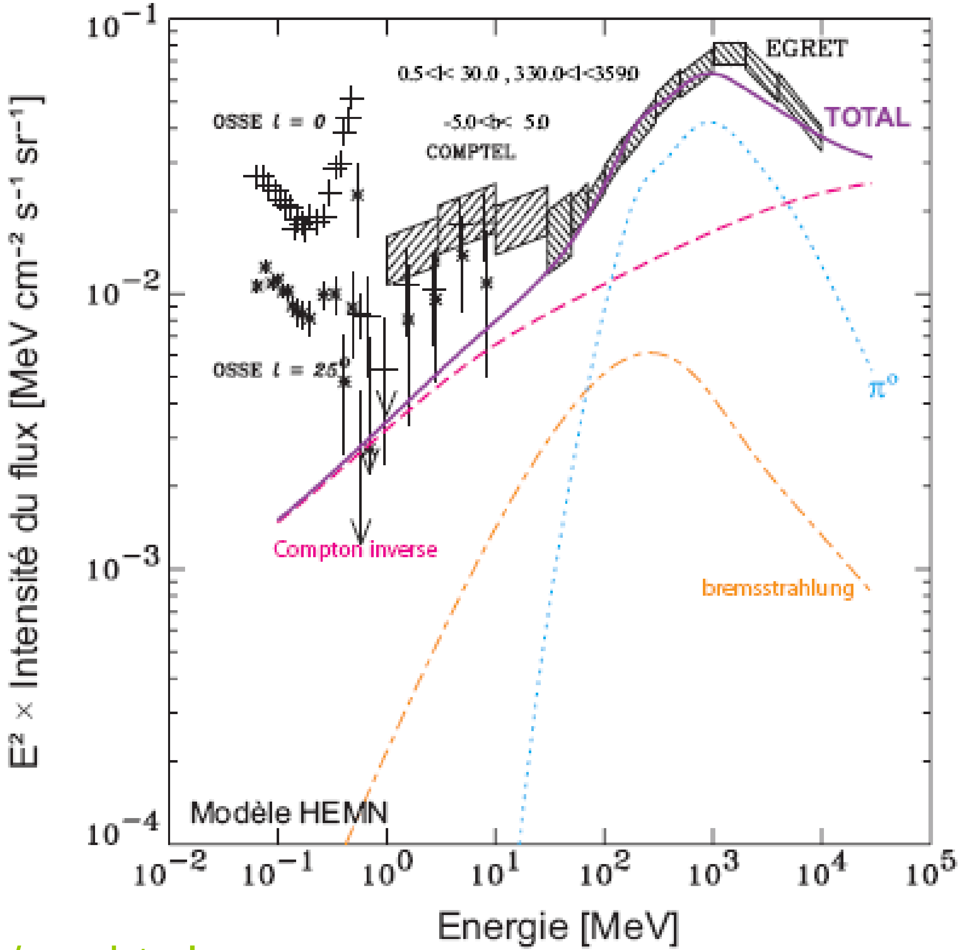






Tracking back the CR flux elsewhere in the Galaxy

- Refined predictions based on detailed simulations
  (magnetic model of the galaxy, CR diffusion equation, matter density maps…):
  for ex: GALPROP program

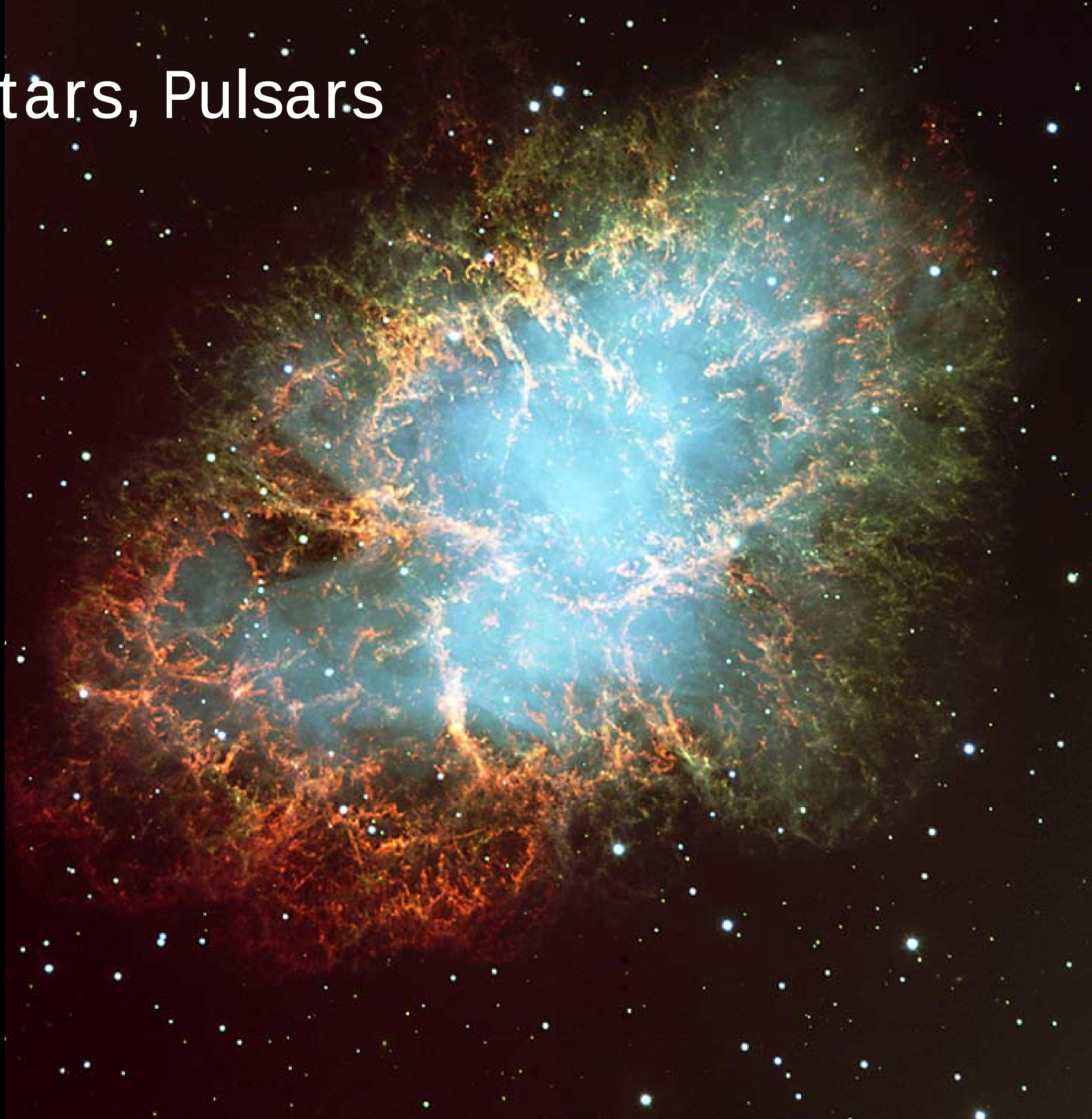


http://www.gamma.mpe-garching.mpg.de/~aws/aws.html

COMPACT OBJECT ENVIRONMENT : NEUTRON STARS AND PULSARS BLACK HOLES

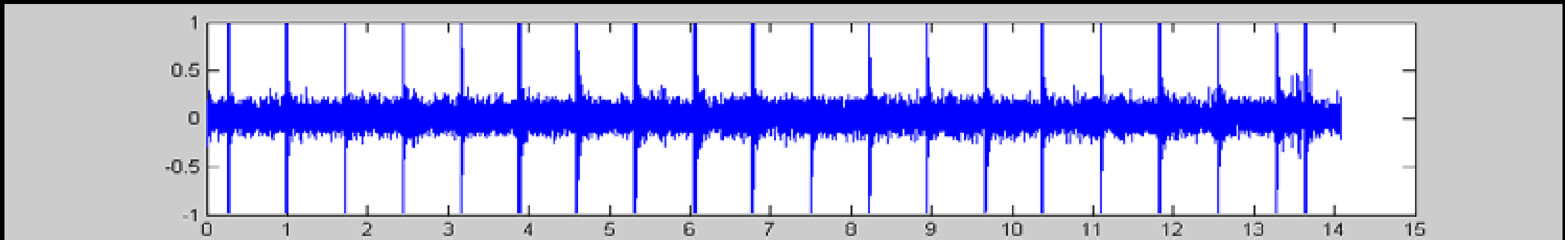
Neutron stars, Pulsars

Journey into the
Crab nebula.

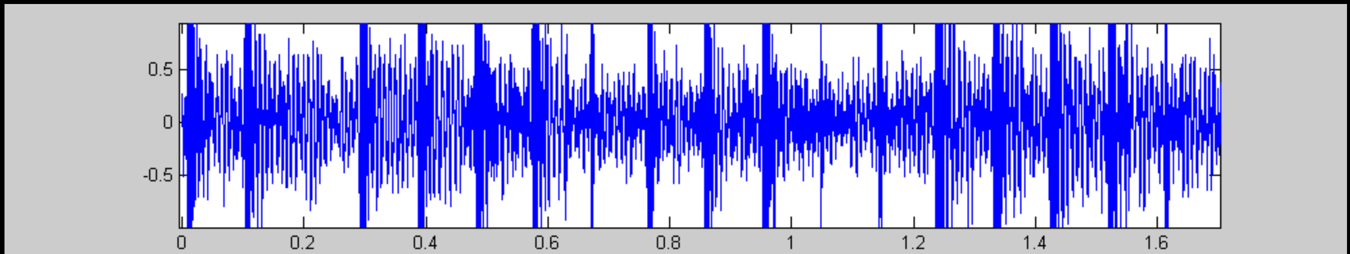


Here is radio pulsar...

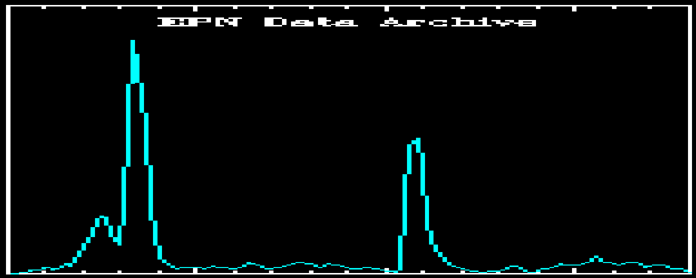
PSR0329+54
~1,40 Hz



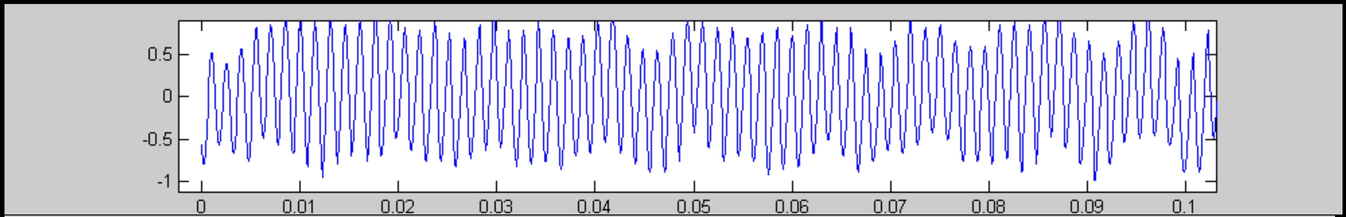
Vela
PSR0833-45
~11Hz

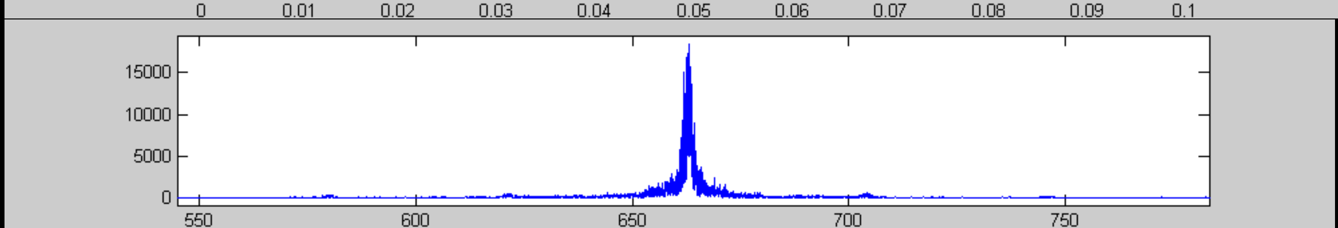


Crabe
PSR0531+21
~30 Hz



PSR1937+21
~641 Hz





Neutron stars

- Ultra compact objects:
  - Density ~ nuclear matter (10^{13} ´ water !)
  - sphere ~ 10 km radius
  - gravitation at surface ~ 5´ 10^{10} ´ gravitation on earth
- Density and structure varies with :

Magnetic field of a neutron star

- Most celestial bodies bear a magnetic field (i.e. the sun: 10^{-3}Tesla)
- A the surface of a neutron star, it is $\approx$ 10 million Tesla
- This value is intrinsically linked to the rapid rotation (conservation of rotation kinetic energy and magnetic energy $\rightarrow$ concentration)
- A kind of giant COSMIC DYNAMO : rotation of $\vec{B} \Rightarrow \vec{E} \Rightarrow 10^{18} V$!!
- The electric force at the surface is $\gg$ gravitationnal force !
  $\rightarrow$ charged particles are expelled and accelerated
  $\Rightarrow 10^{38}\ e^-$ per second radiating synchrotron light.
- The strong anisotropy of the radiation is badly understood but probably due to the intense $\vec{E}$ and $\vec{B}$ fields near the "polar caps".

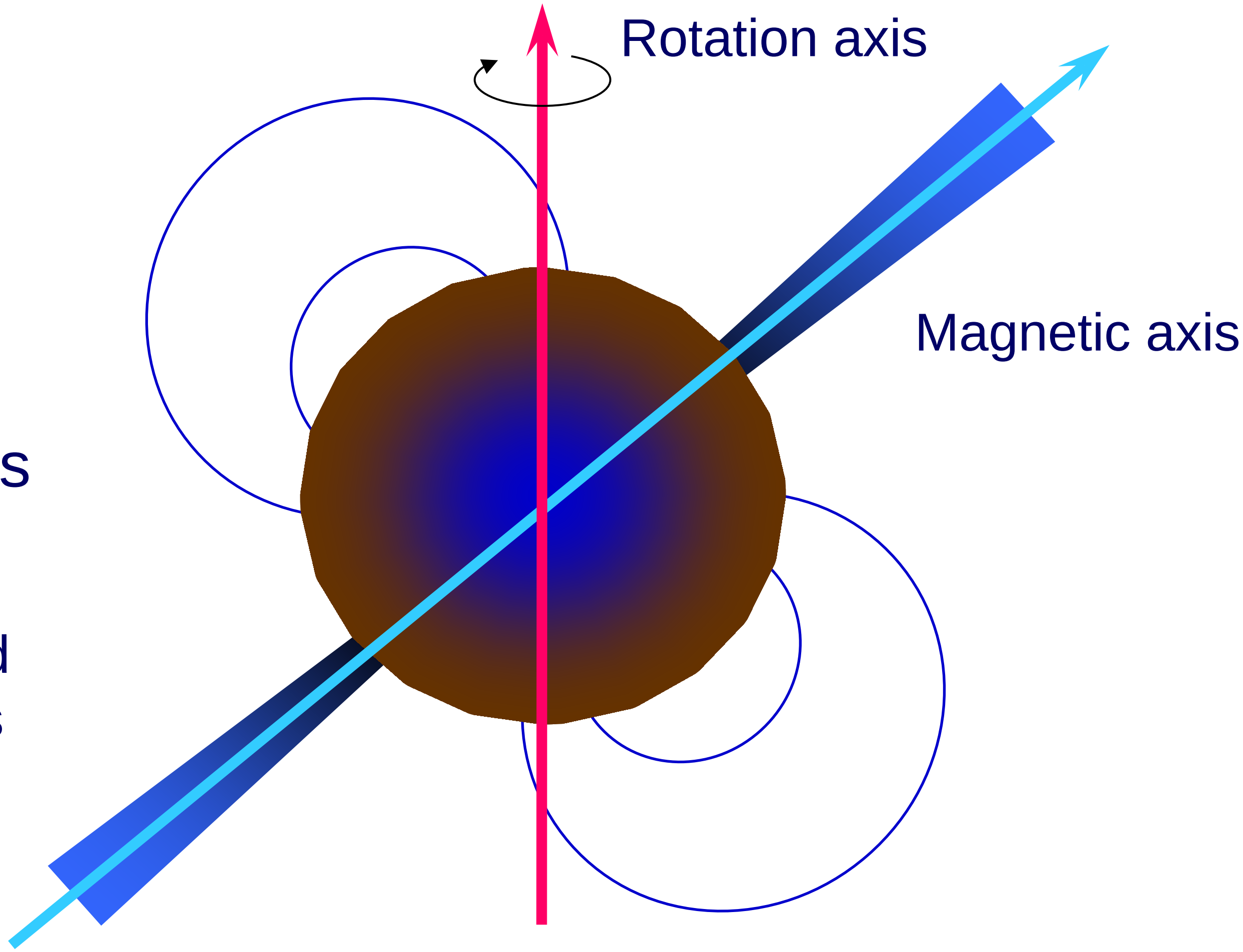
Unipolar Induction

– **Neutron star** :
a rotating magnet with a magnetosphere

Neutron star

Mass = 1.4 M¤
Radius = 10 km
Rotation period = 1 s

Radio beam emitted
along magnetic axis



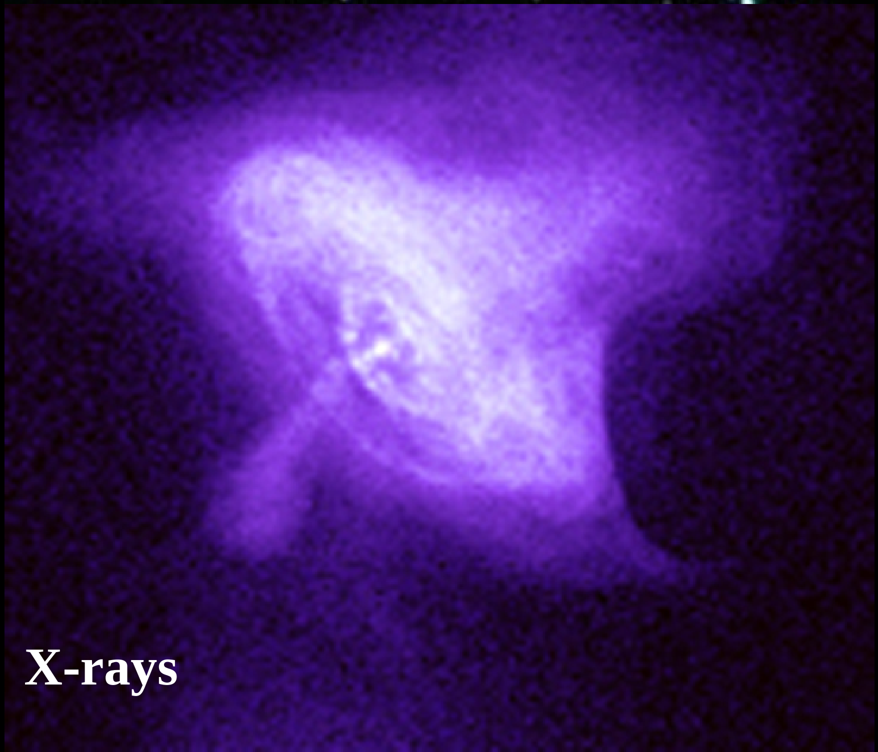
Induction unipolaire

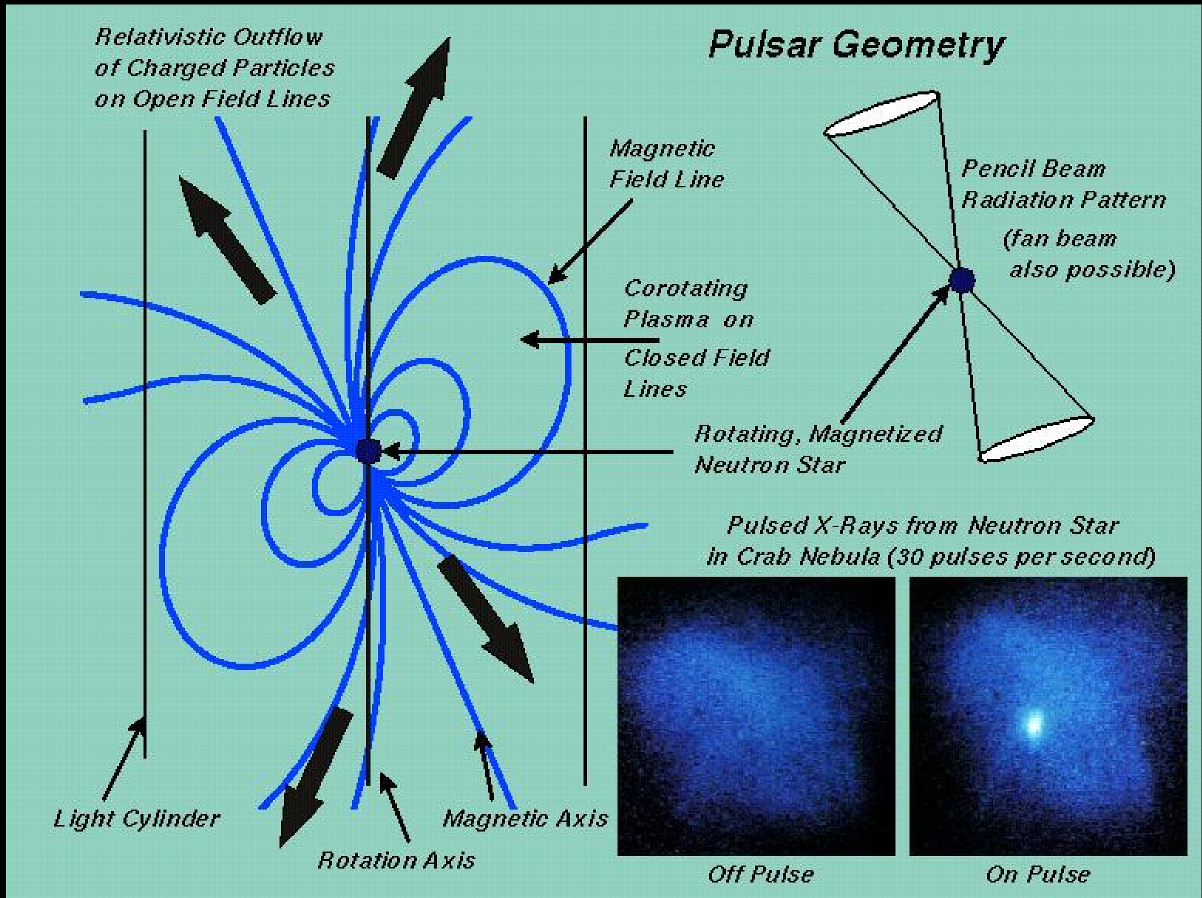
– **Neutron star** :
a rotating magnet with a magnetosphere

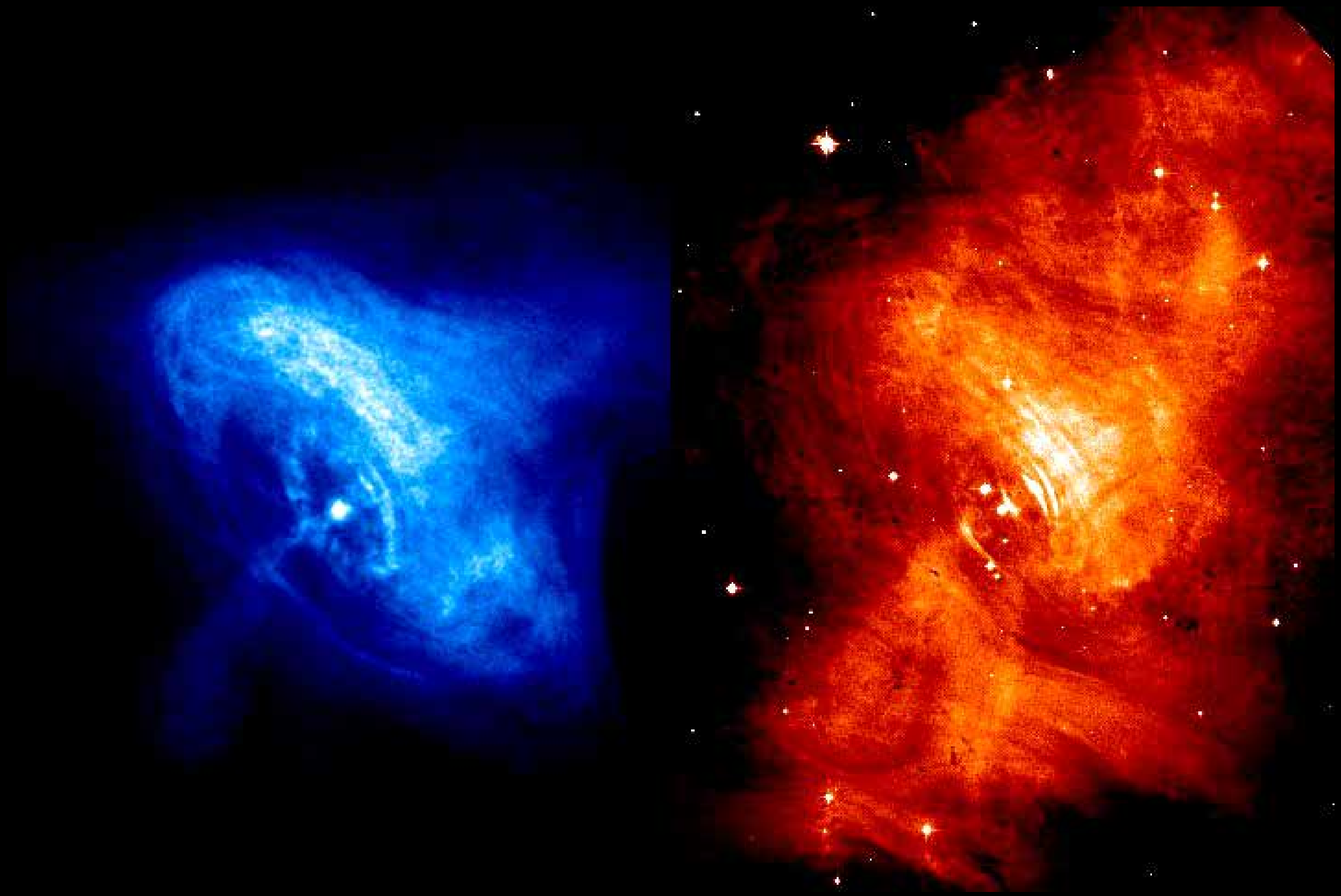
$$\rho_0 = \vec{\nabla}.\left(\frac{\left(\vec{\Omega}\times\vec{r}\right)\times\vec{B}}{4\pi c}\right) = \frac{\vec{\Omega}.\vec{B}}{2\pi c\left(1-\left|\vec{\Omega}\times\vec{r}/c\right|^2\right)}$$

$$\Delta V \approx \frac{\Omega^2 B_s R^3}{c^2} = 3\times 10^{16}\ \Omega_2^2\ B_{13}\ R_6^3\ \text{Volts}$$


Optical

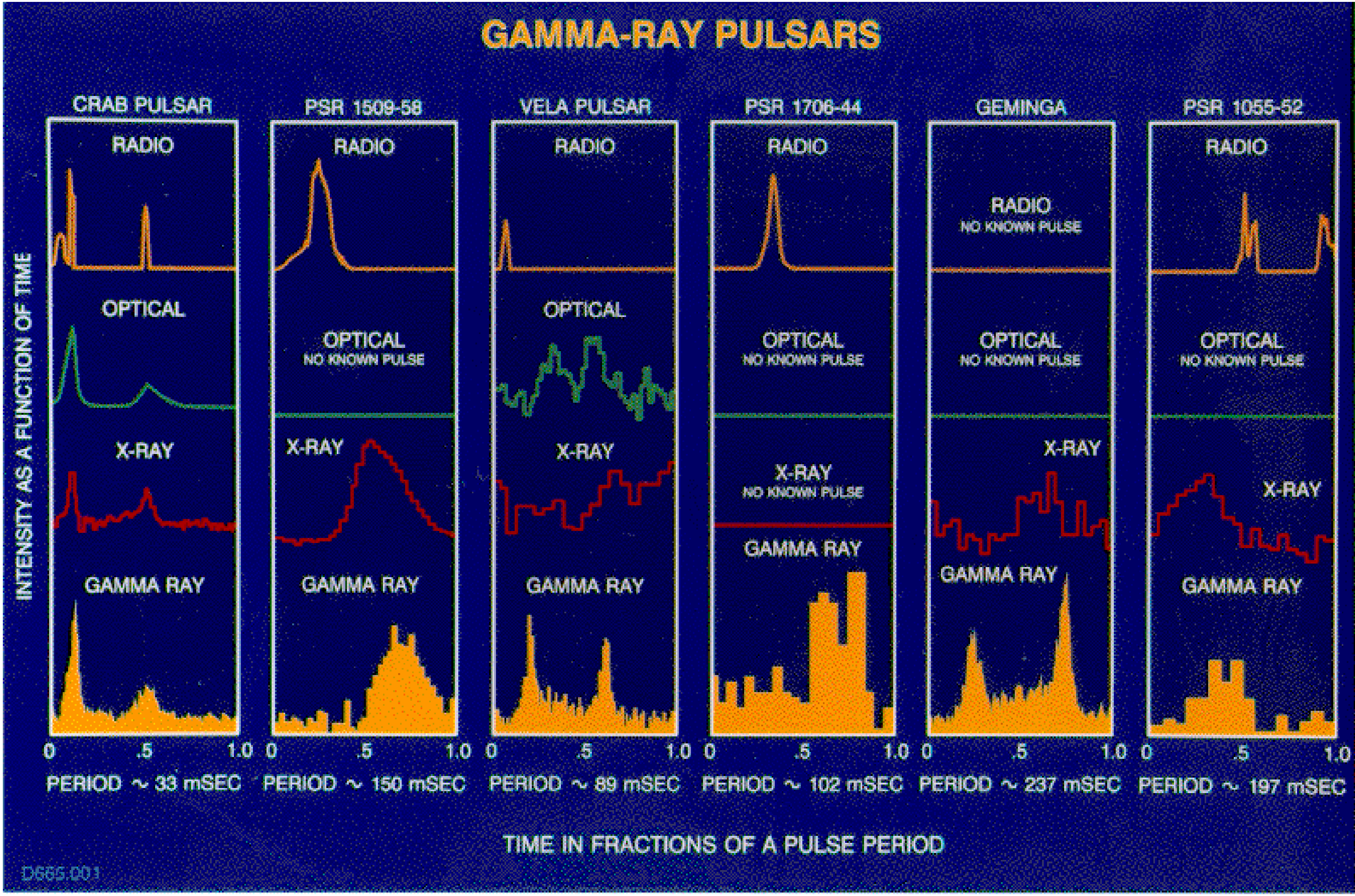

X-rays


Pulsar Geometry
Relativistic Outflow of Charged Particles on Open Field Lines
Magnetic Field Line
Pencil Beam Radiation Pattern
(fan beam also possible)
Corotating Plasma on Closed Field Lines
Rotating, Magnetized Neutron Star
Pulsed X-Rays from Neutron Star in Crab Nebula (30 pulses per second)
Light Cylinder
Rotation Axis
Magnetic Axis
Off Pulse
On Pulse



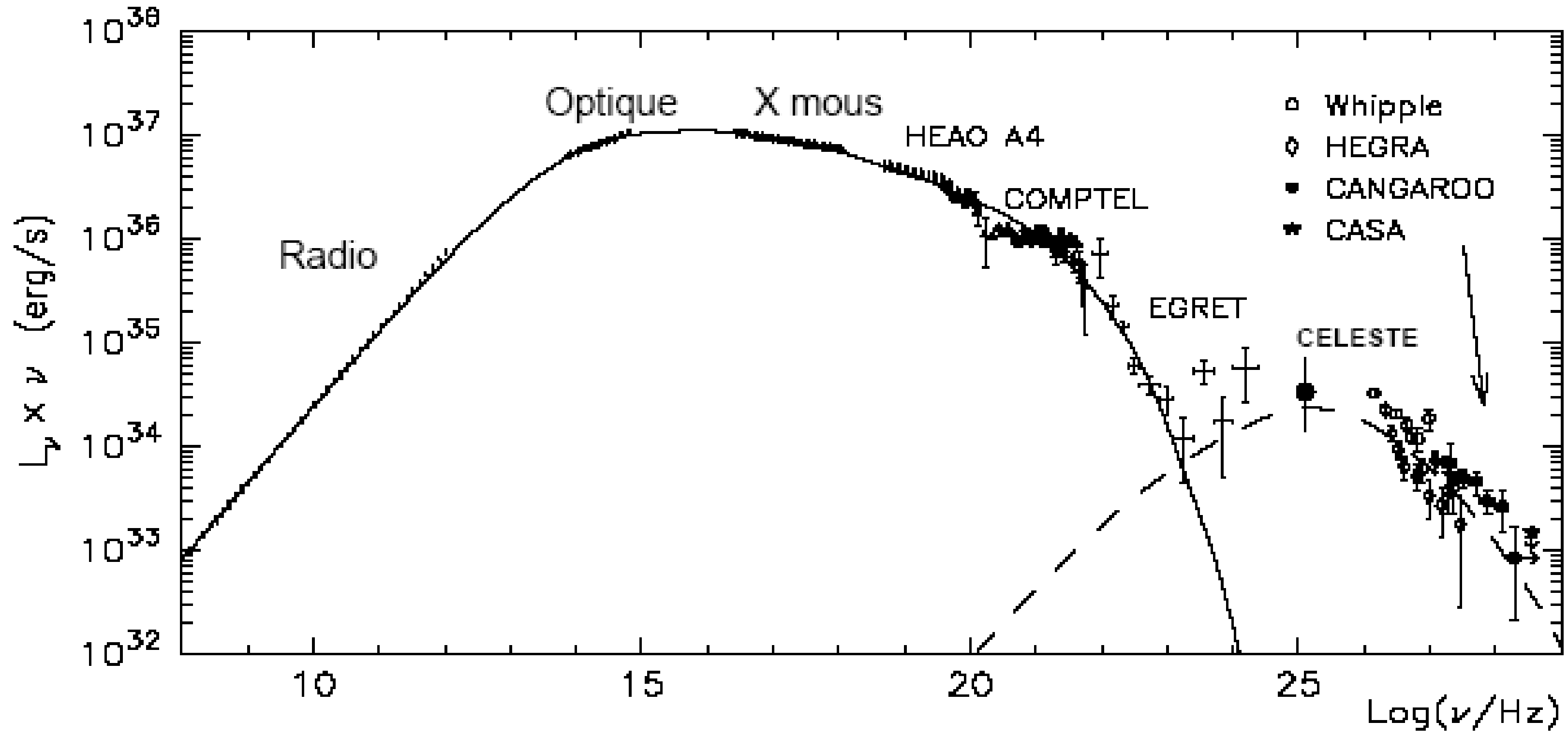
The Crab X-source, the central engine, X (Chandra) et optical (HST)

7 images from november 2000 to april 2001, showing the time and space coincidence between the shocks observed at the very center of the system (close to the pulsar) and the X-ray radiation produced by the accelerated VHE electrons


GAMMA-RAY PULSARS
CRAB PULSAR
PSR 1509-58
VELA PULSAR
PSR 1706-44
GEMINGA
PSR 1055-52
RADIO
OPTICAL
X-RAY
GAMMA RAY
NO KNOWN PULSE
INTENSITY AS A FUNCTION OF TIME
0
.5
1.0
PERIOD ~ 33 mSEC
PERIOD ~ 150 mSEC
PERIOD ~ 89 mSEC
PERIOD ~ 102 mSEC
PERIOD ~ 237 mSEC
PERIOD ~ 197 mSEC
TIME IN FRACTIONS OF A PULSE PERIOD
D665.001

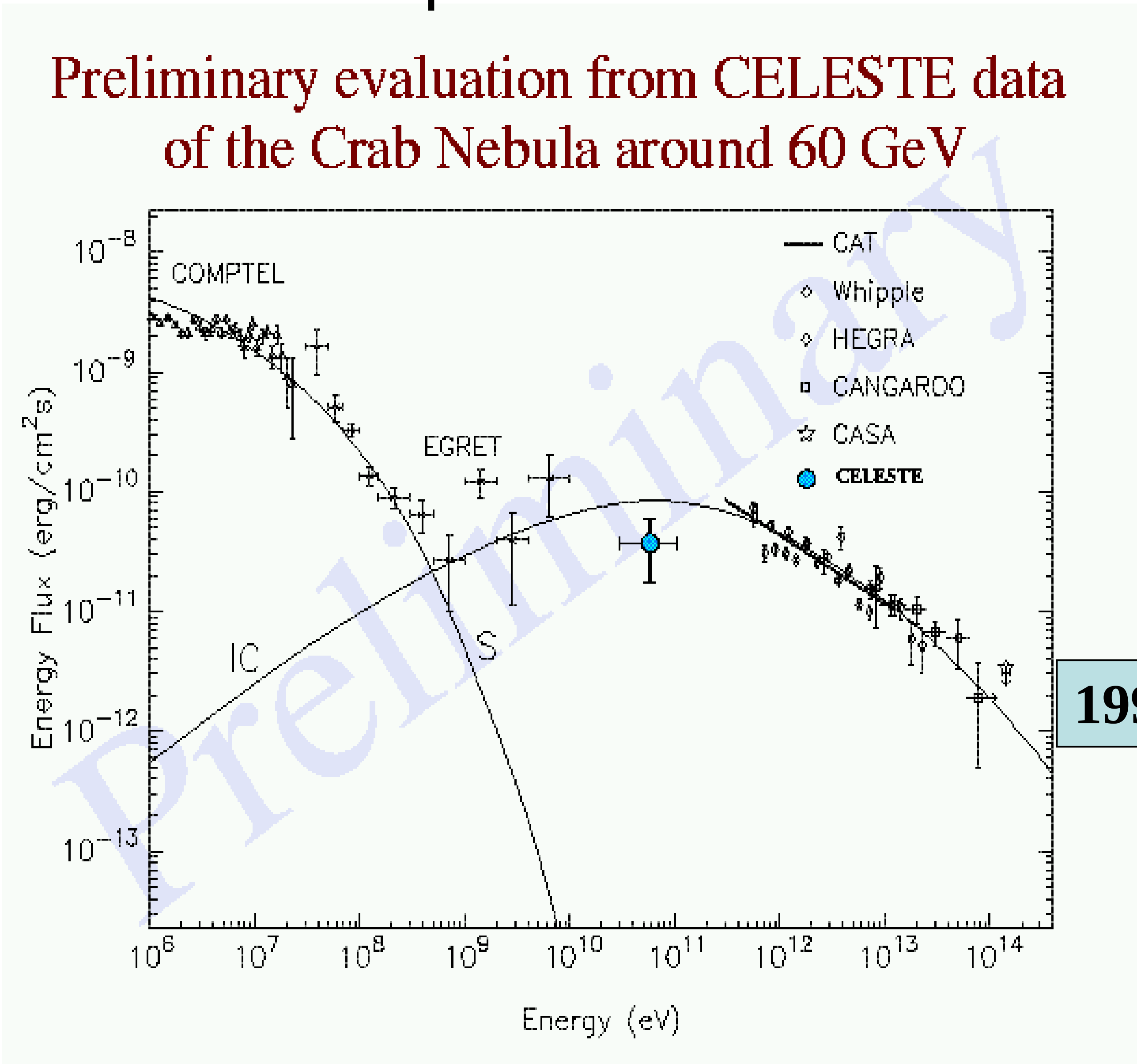
The Crab Nebula
a case study for VHE gamma-rays VHE

- Steady emission observed at all wavelengths.
- First point like gsource identified
- Intense flux: a "standard candle" for °-ray observatories



The Crab Nebula
a case study for VHE gamma-rays VHE

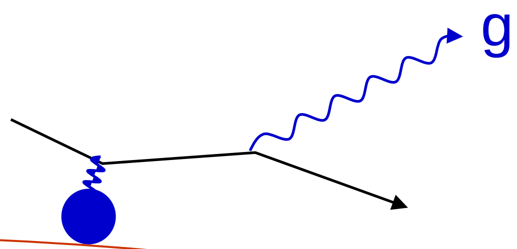
- Two distinct componants



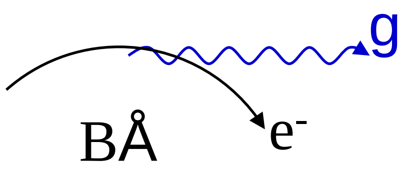
1999-2000

VHE photons sources

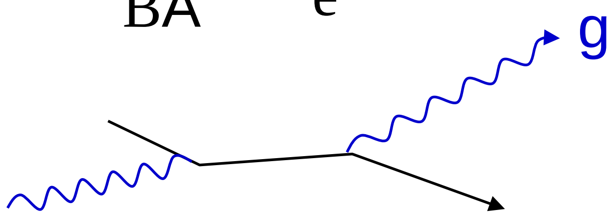
- °-ray production:
  - Electro-magnetic processes :
    - Bremsstrahlung



    - Synchrotron radiation



    - Inverse Compton scattering



  - Hadronic processes :

$$(p, N) + (N\gamma) \rightarrow \pi_0 + \dots$$
$$\hookrightarrow \gamma\gamma$$

The Crab Nebula
a case study for VHE gamma-rays VHE

- TeV emission is based on Inverse Compton scattering

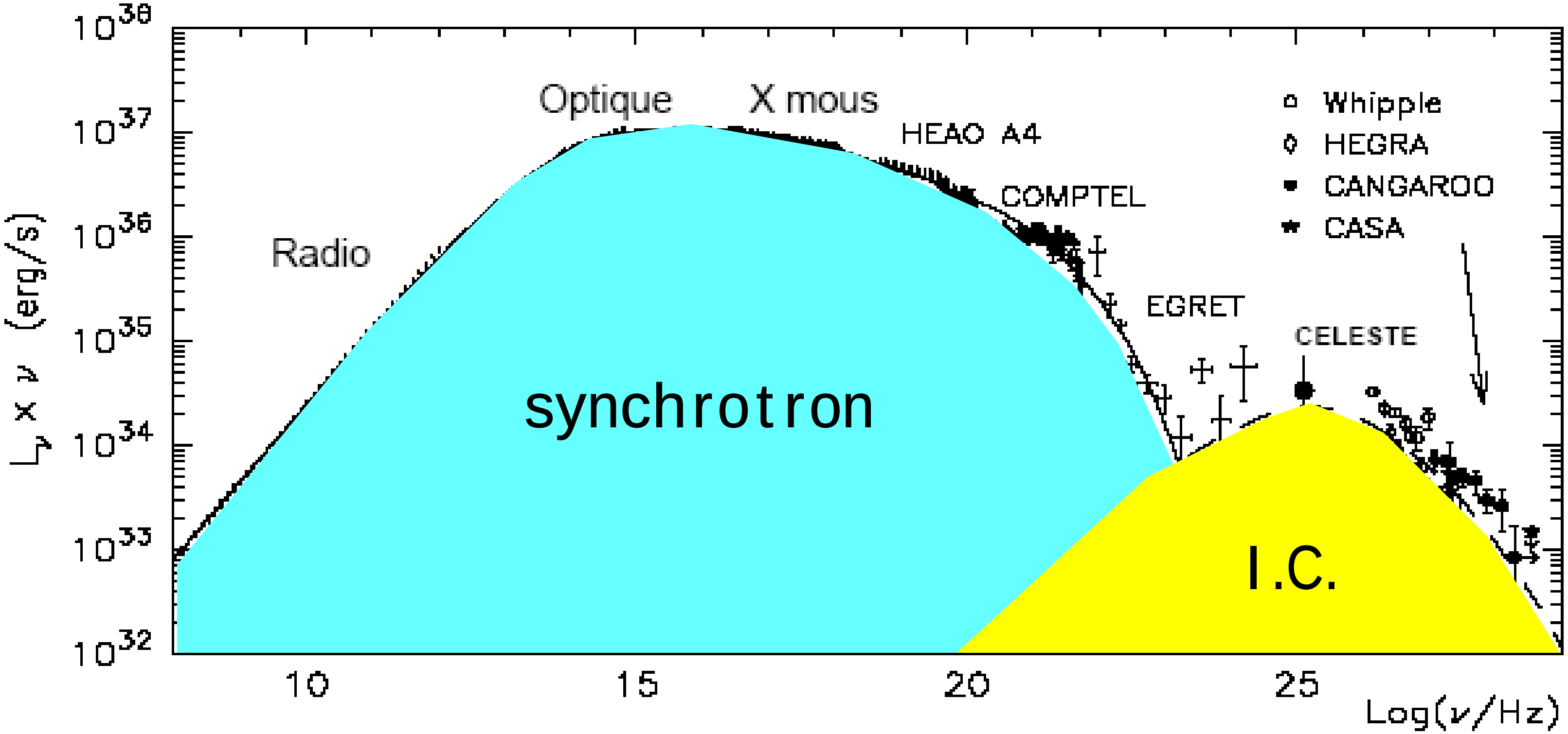
- The e^- producing synchrotron radiation can interact with these synchrotron photons.

- If $\lambda_{\text{photon cible}} \ll \lambda_{\text{compton}} = \frac{h}{m_e c}$ it is possible to transfer most of the incident e^- energy to the photon. ex: A $10^{13} eV$ e^- can boost an I.R. photon I.R. into a VHE gamma-ray.

- These "Self Synchrotron Compton" models reproduce both GeV and TeV spectra.

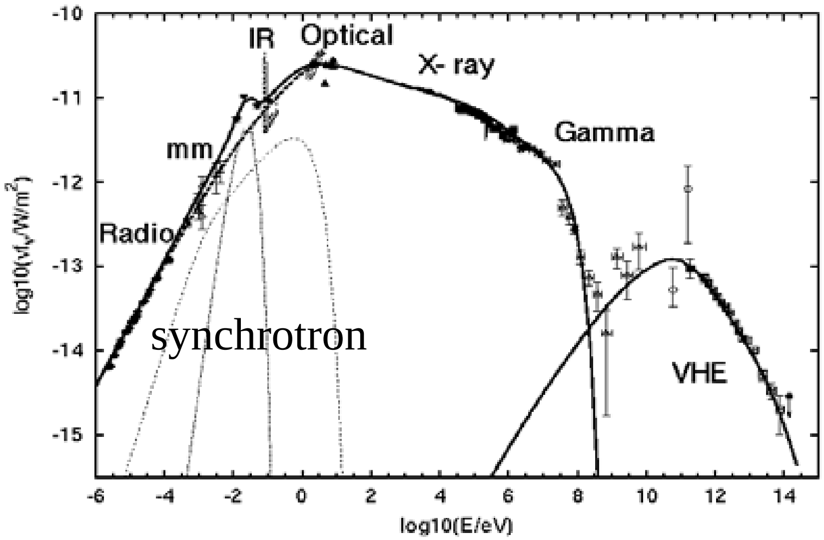
La Nébuleuse du Crabe et les rayons gamma VHE (~TeV)

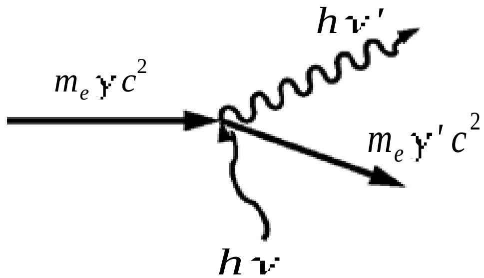
Model SSC : Synchroton Self Compton



PWN emission mechanisms: the Crab Nebula

- Assume leptonic model: synchrotron and Inverse Compton emission
- Relativistic electrons and positrons created and accelerated by the pulsar

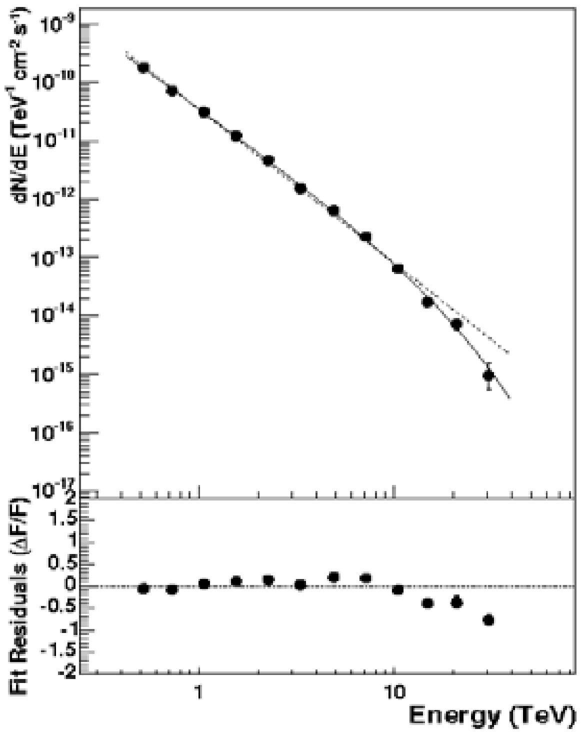




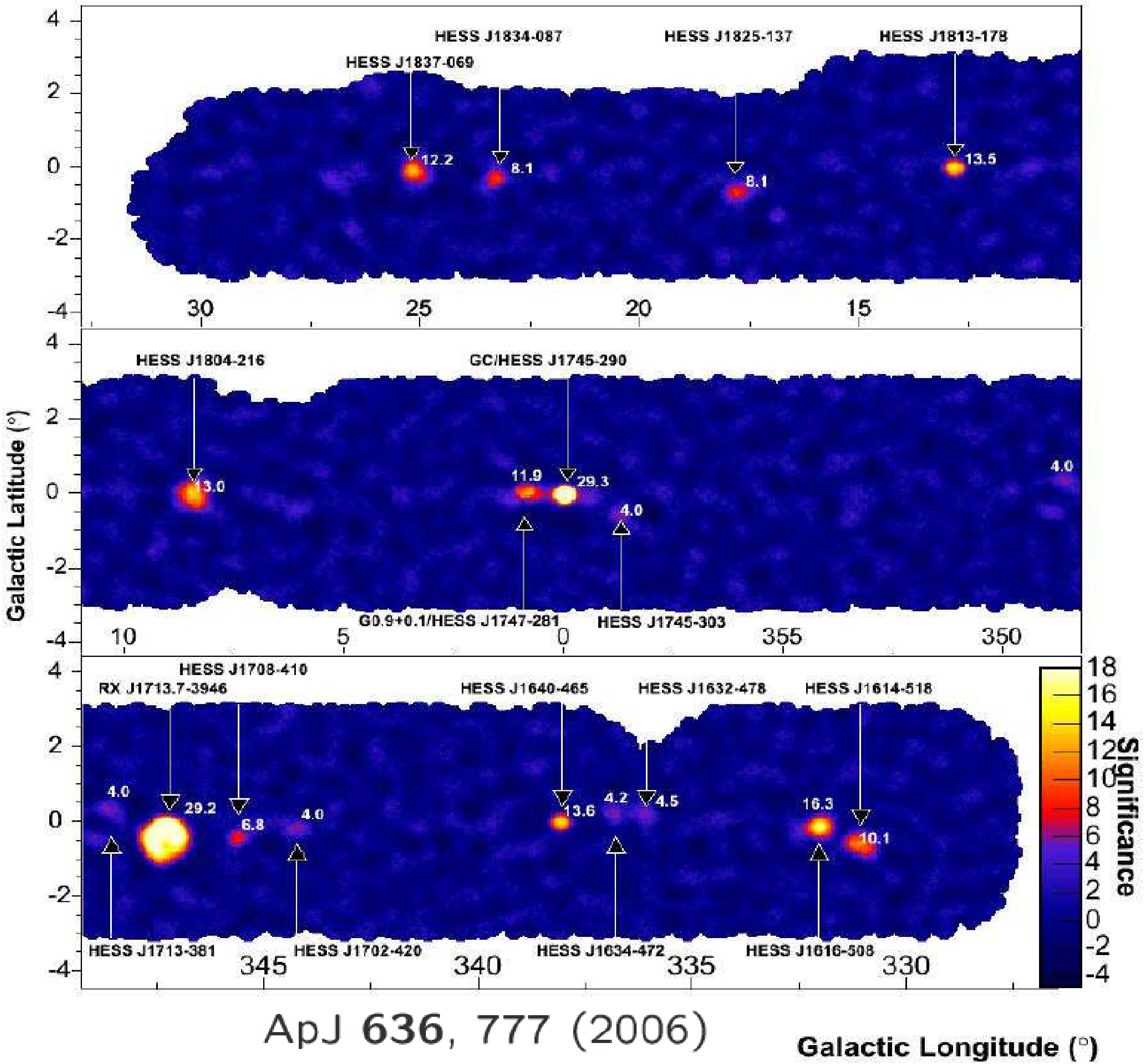
Target photons : CMB, interstellar IR, stellar photons, synchrotron (SSC)…

Radio, *optical*, *X-rays*

H.E.S.S. spectrum (*A&A* 2006 in press, astro-ph/0607333): Spectral curvature, Consistent with IC expectations



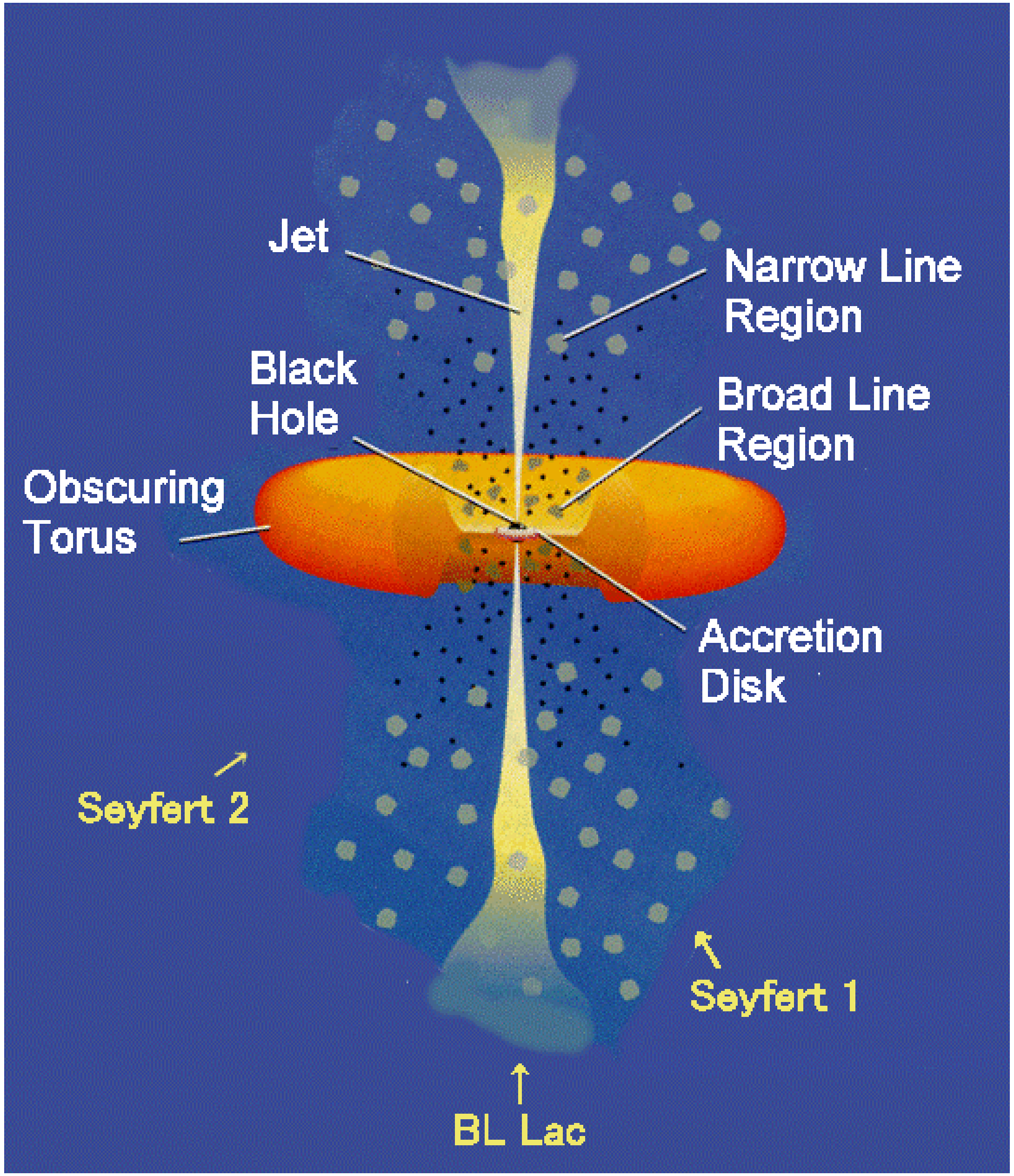
Many recently discovered sources in the Galactic plan by HESS, MAGIC and VERITAS large angle surveys



ApJ **636**, 777 (2006)

THE ENVIRONMENT OF SUPERMASSIVE BLACK HOLES: ACTIVE GALATIC NUCLEI

AGN: a unified scheme



Active Galactic Nuclei

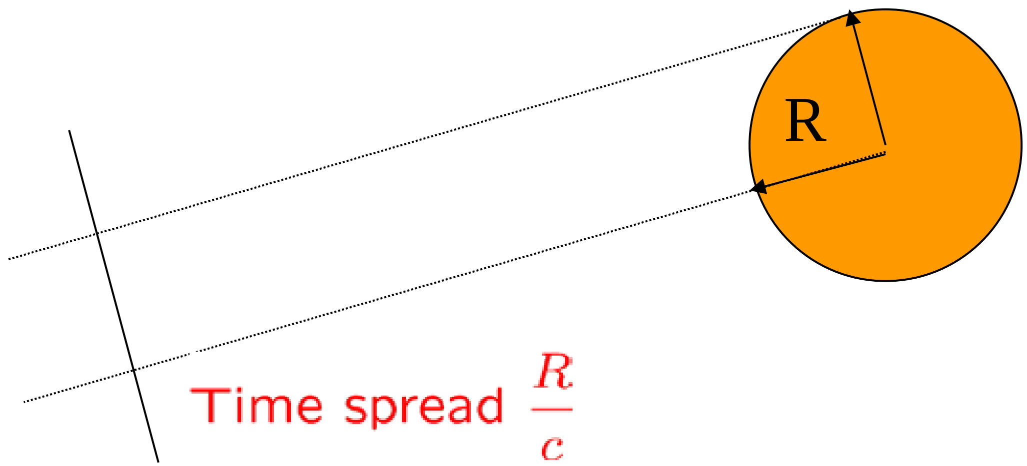
- Galaxies with an important activity of its nucleus
- Numerous classes of objects:
  - Radio-galaxies
  - Radio-quasars
  - BL Lac objects
  - Violently Variables objects (VVO)
  - Radio Quiet Quasars
  - Seyfert Galaxies of type 1
  - Seyfert Galaxies of type 2
  - Low Ionization Nuclear Emission-Line Regions (LINERs)
  - Nuclear HII Regions
  - "Star Burst" Galaxies
- Arbitrary categories et badly mal defined.

What is an AGN

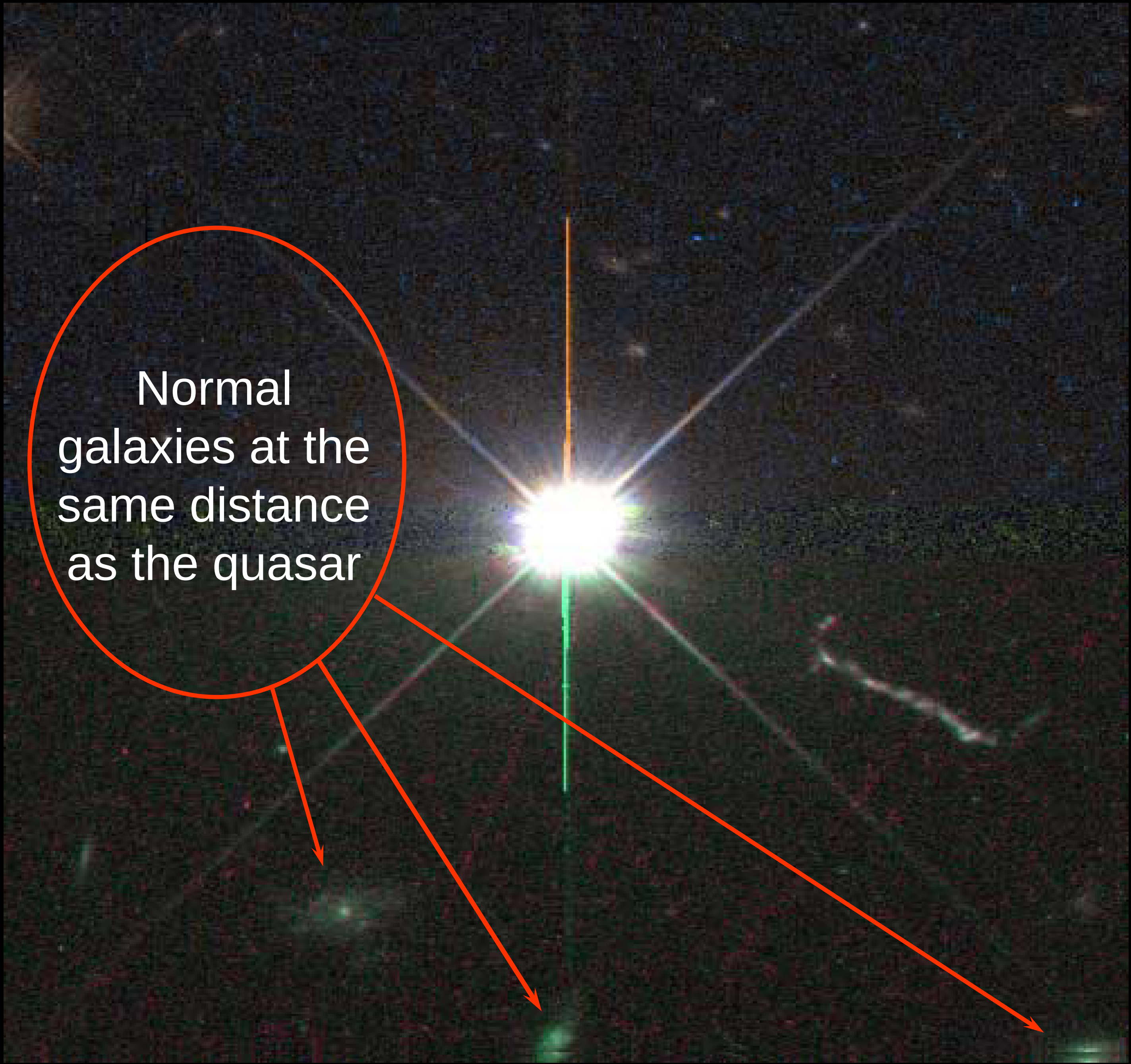
- Elementary AGN model:

  Discovered in the early 60ties when trying to identify the sources in the 3rd Cambridge Catalogue.

  - Superluminous objects
    - Spectral redshift ® cosmological distances
    - For 3C273 ® P ¼ 10^{40} Watts ¼ 10000 the Milky Way ¼ 1000 SN/year
  - Very small objects
    - Variability criteria ® d ¼ solar system size !!



The Quasar 3C 273



In 1963, 3C 273, the first quasar, was discovered. They can be up to 1000 more luminous than the massive host galaxy. The emission originates in the active nucleus of the galaxy.

What is an AGN

- Elementary AGN model:

  Discovered in the early 60ties when trying to identify the sources in the 3rd Cambridge Catalogue.

  – Objects from the past
    - The most distance quasars observed (z ¼ 5) are at distances of the order of the radius of the observable universe!
    - No such objects in our close environment ® must be linked to a phase of the galaxy evolution.

What is an AGN

- Necessarily a gravitational engine

Thermonuclear falls short by a factor 10 to 100 !

Form of energy release	Efficiency of energy production (wrt mc^2)
• Chemical energy	10^{-9}
• Nuclear energy	10^{-2}
• Accretion of mass onto non-rotating black holes	$6x10^{-2}$
• Accretion onto maximally rotating Kerr black holes	0.42
• Rotational energy of maximally rotating Kerr black holes	0.29

What is an AGN

- Necessarily a gravitational engine

$$E_{tot} = 0 \Rightarrow E_{cin} = \frac{GmM}{R} \Rightarrow \text{R must be small}$$

Huge mass, small volume $\Rightarrow$ BLACK HOLE !!!

With $R_S = \frac{2GM}{c^2} \Rightarrow E_{rad.max} \approx \frac{GM}{R_S}\frac{R_S}{R}m = \frac{1}{2}\frac{R_S}{R}mc^2$

Accounting for relativistic corrections and ergosphere rotation
$\rightarrow$ conversion efficiency $m \rightarrow E$ few 10%

What is an AGN

- Necessarily a gravitational engine

One gets 10^{40} Watts with $15 M_\odot$/year.
Over 10^8 years (age measured from radio lobes) $\Rightarrow$ $10^9 M_\odot$
$\Rightarrow R_S \approx 10^{13} m$ compatible with variability.

More over, one is below Eddington luminosity
(radiation pressure = gravitational forces)

$$\left(L \approx \frac{4\pi c}{\sigma_T} G M m_H\right) \text{ de } \left(\frac{\sigma_T L}{4\pi R^2} = \frac{G M m_H}{R^2}\right)$$

$\Rightarrow$ PLAUSIBLE.

What is an AGN

- Accretion of matter:

  - In absence of angular momentum:
    $\vec{v}_\infty$ directed toward the BH: $\frac{GMm_H}{R} = kT$
    $\Rightarrow$ almost all is radiated in X or γ rays
    INCOMPATIBLE with MEASUREMENTS

  - With angular momentum:
    Intense radiation in the UV : OBSERVED

    $\Rightarrow$ Anisotropic Accretion
    in the shape of an ACCRESTION DISK

What is an AGN

- Jets:

  Relativistic particles constituting the jets are accelerated when reflecting on the bondaries of a magnetic "tunnel" induced by the rotating plasmas near the BH.

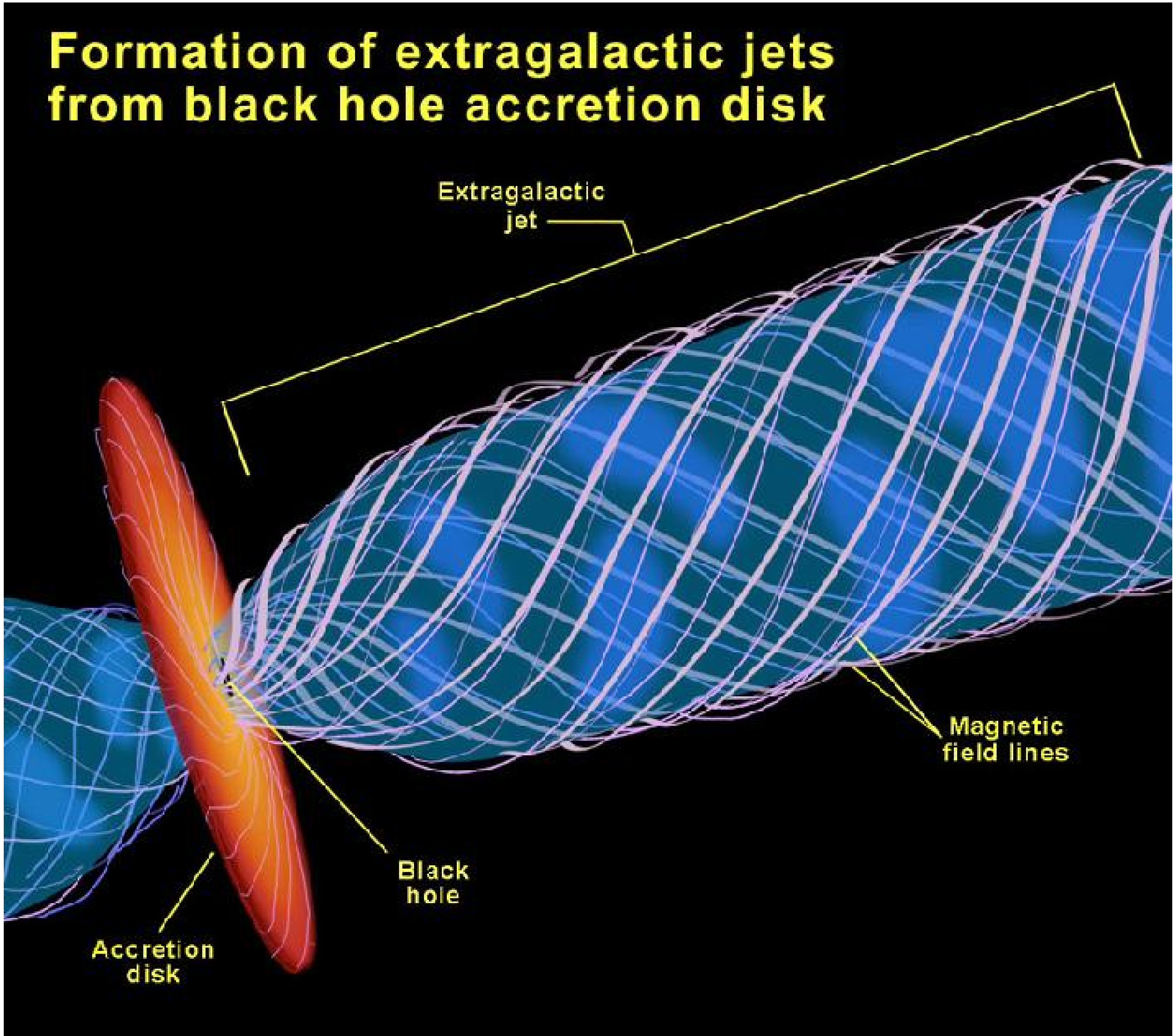
- Hot spots

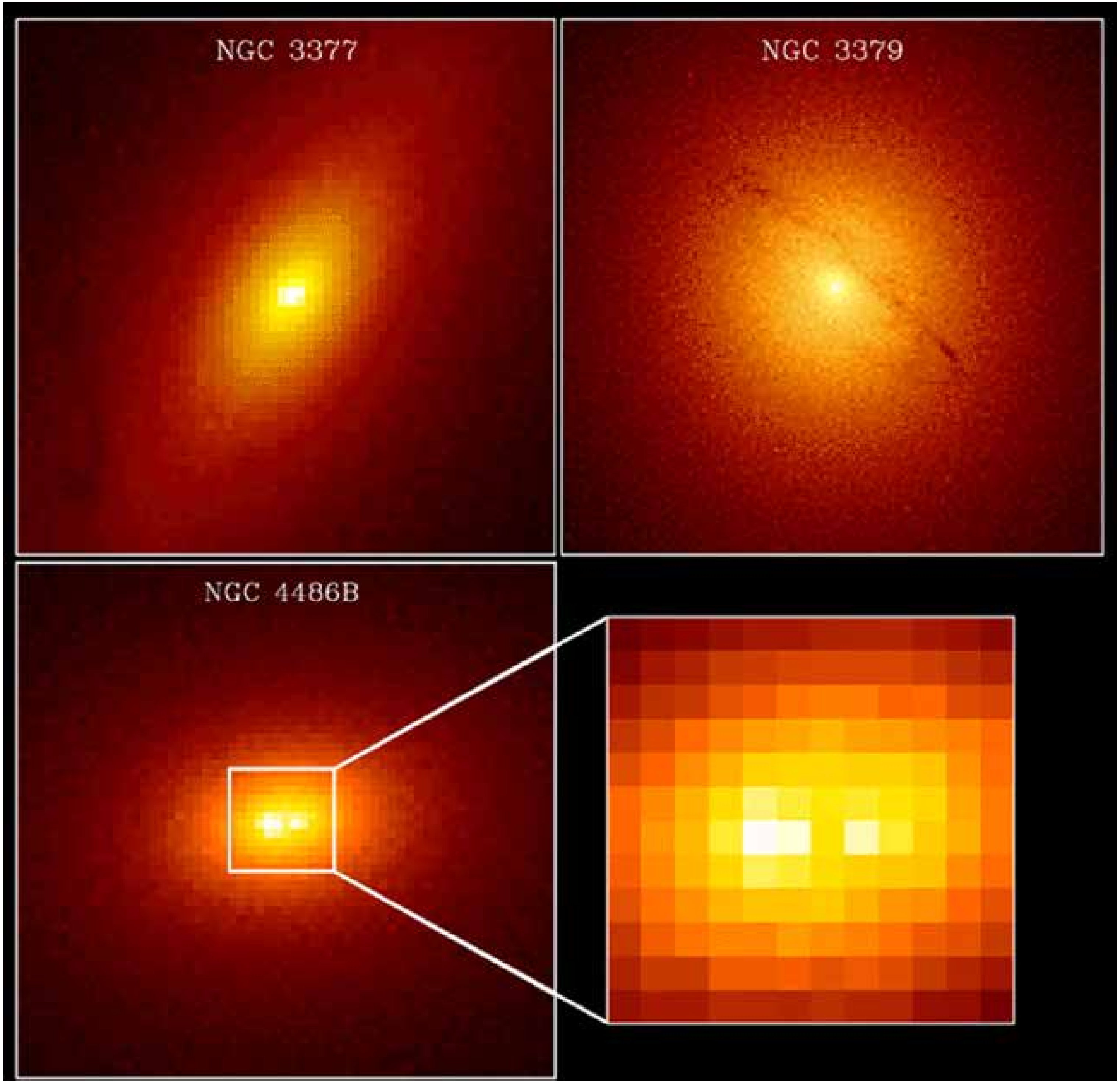
Terminal shocks (jet-IGM): large size moderate field regions
$\Rightarrow$ plausible place for UHECR production.

$\rightarrow$ explains many observed phenomena with both leptonic and hadronic models.

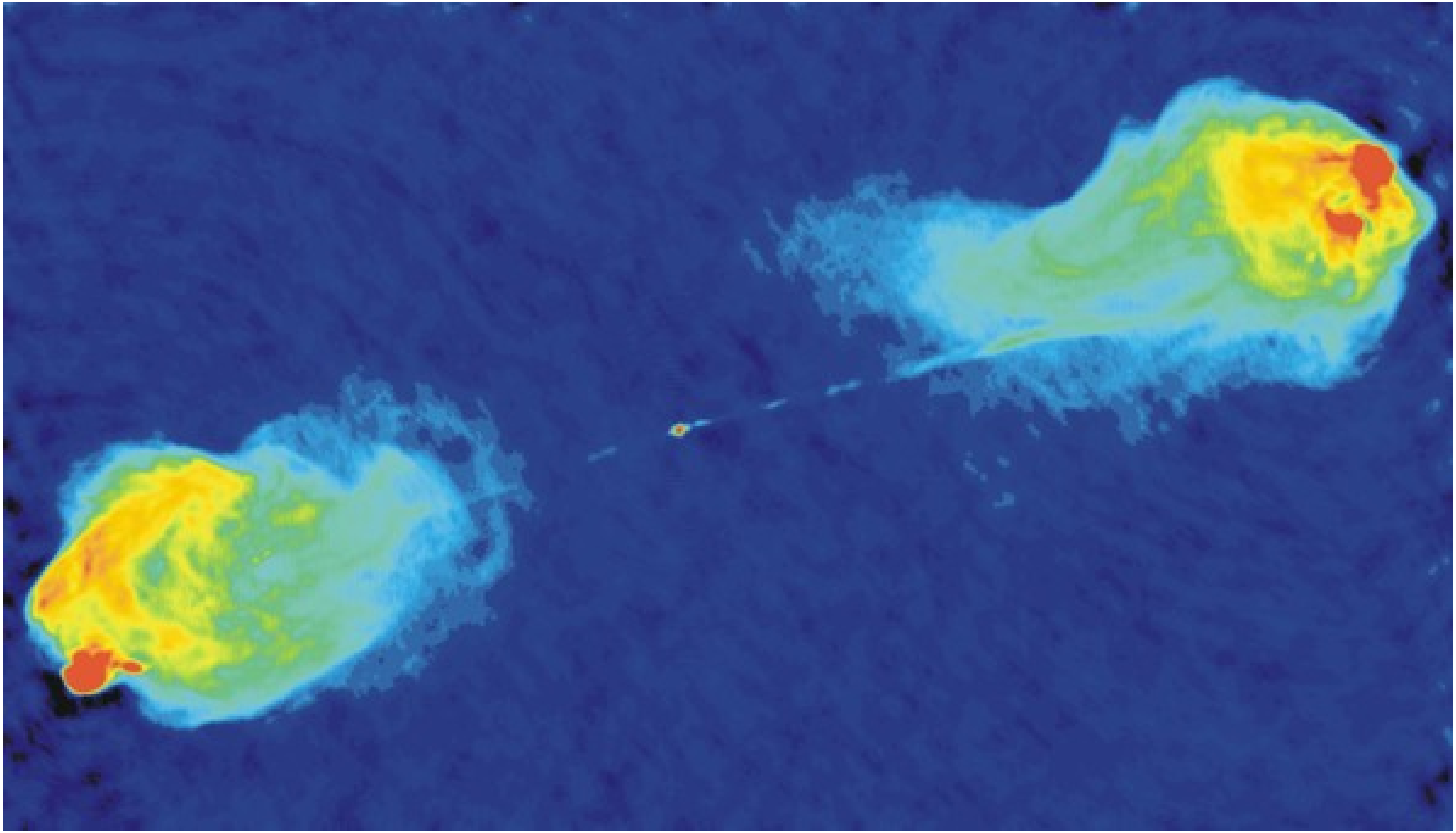
Qu'est-ce qu'un AGN

- Jets :

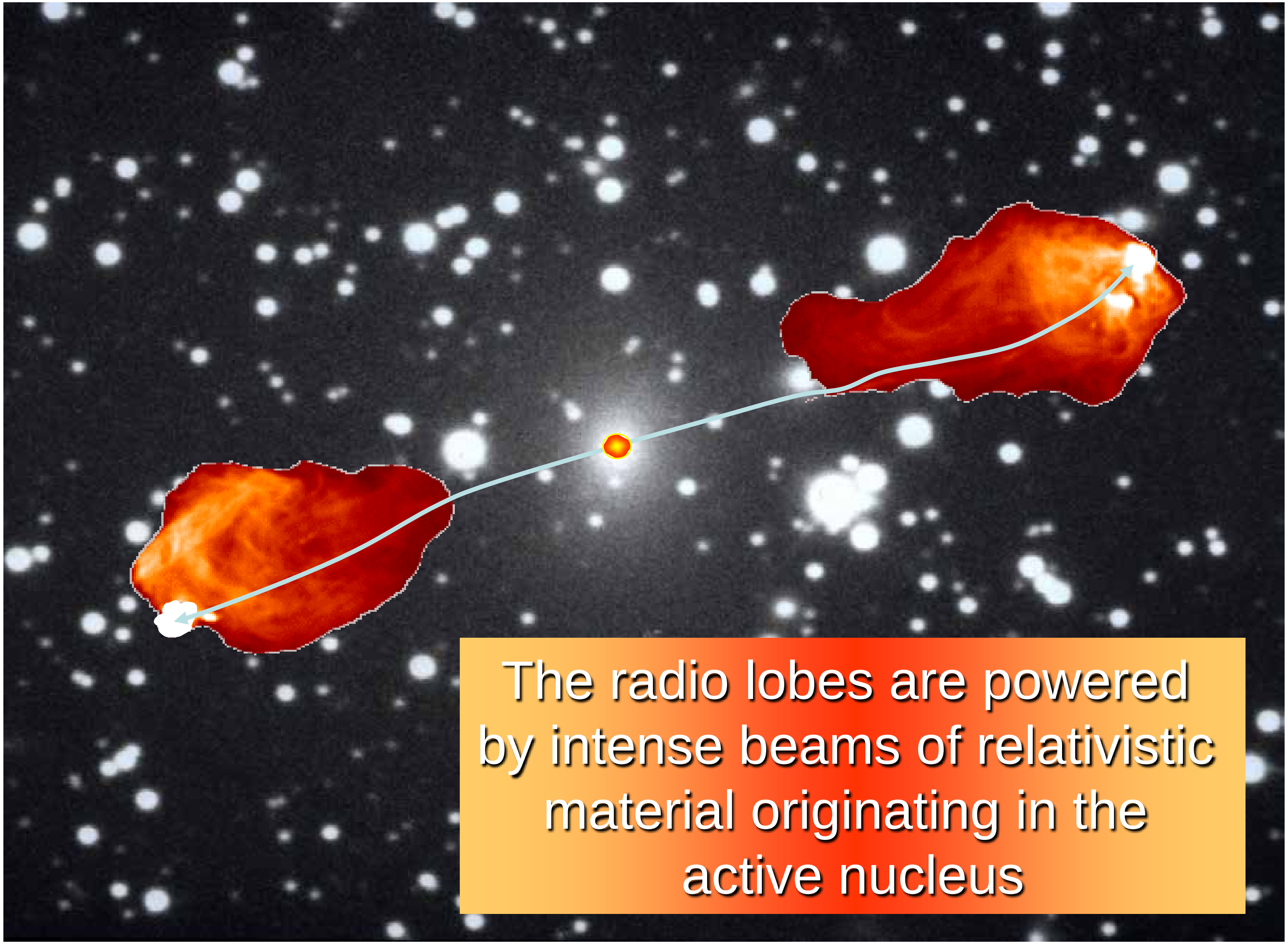



NGC 3377
NGC 3379
NGC 4486B

Cygnus A

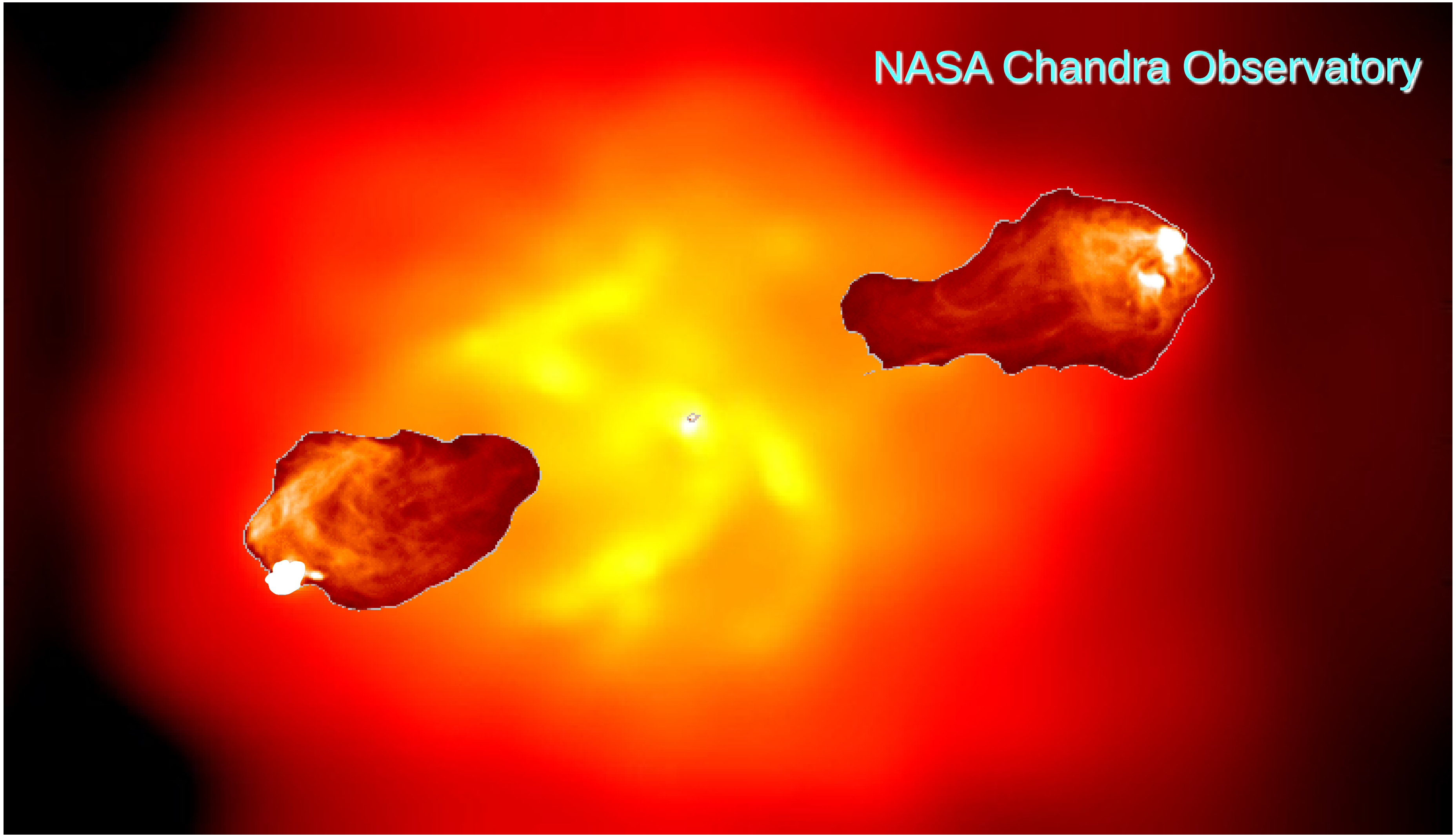


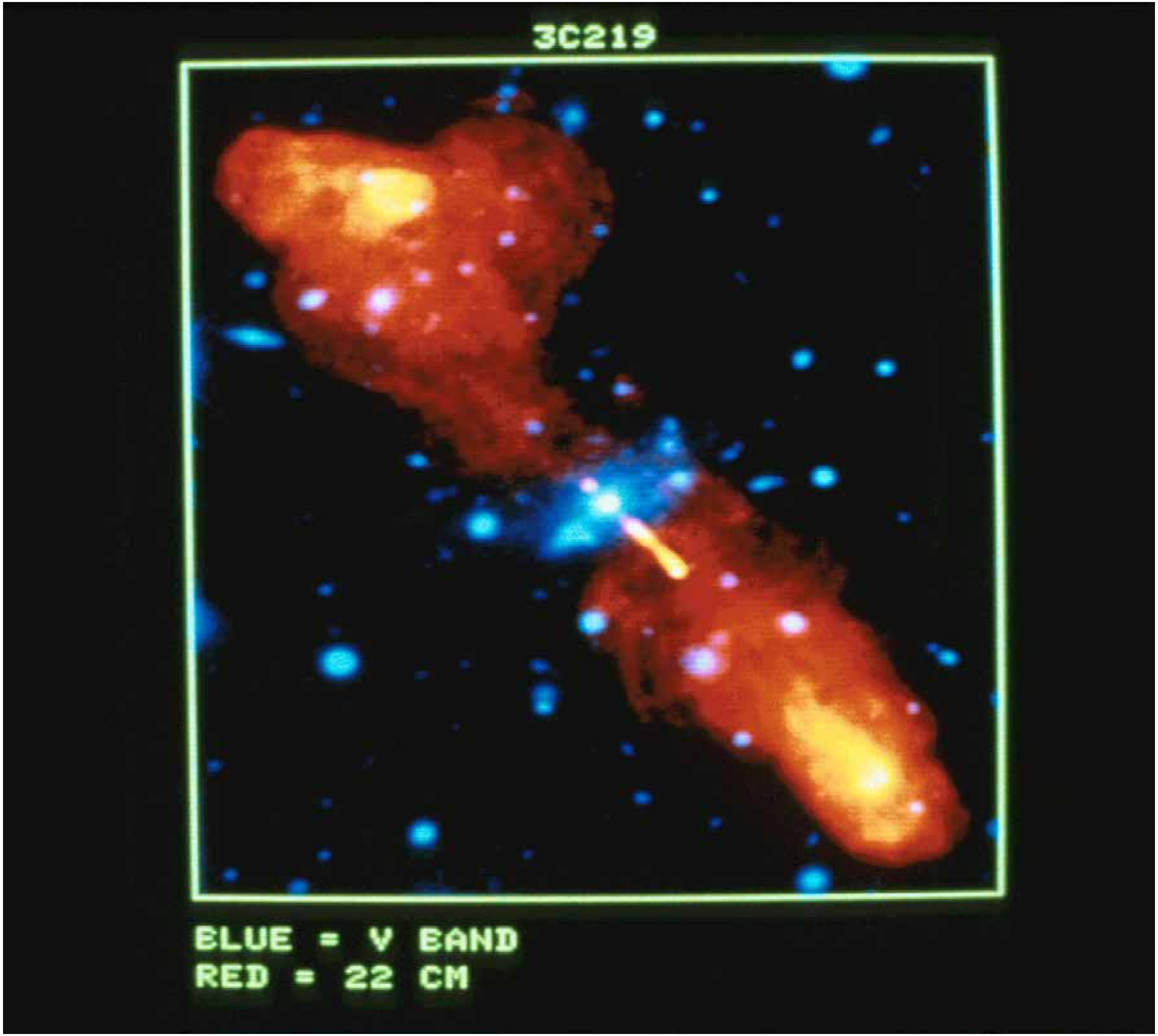
The Radio Galaxy Cygnus A

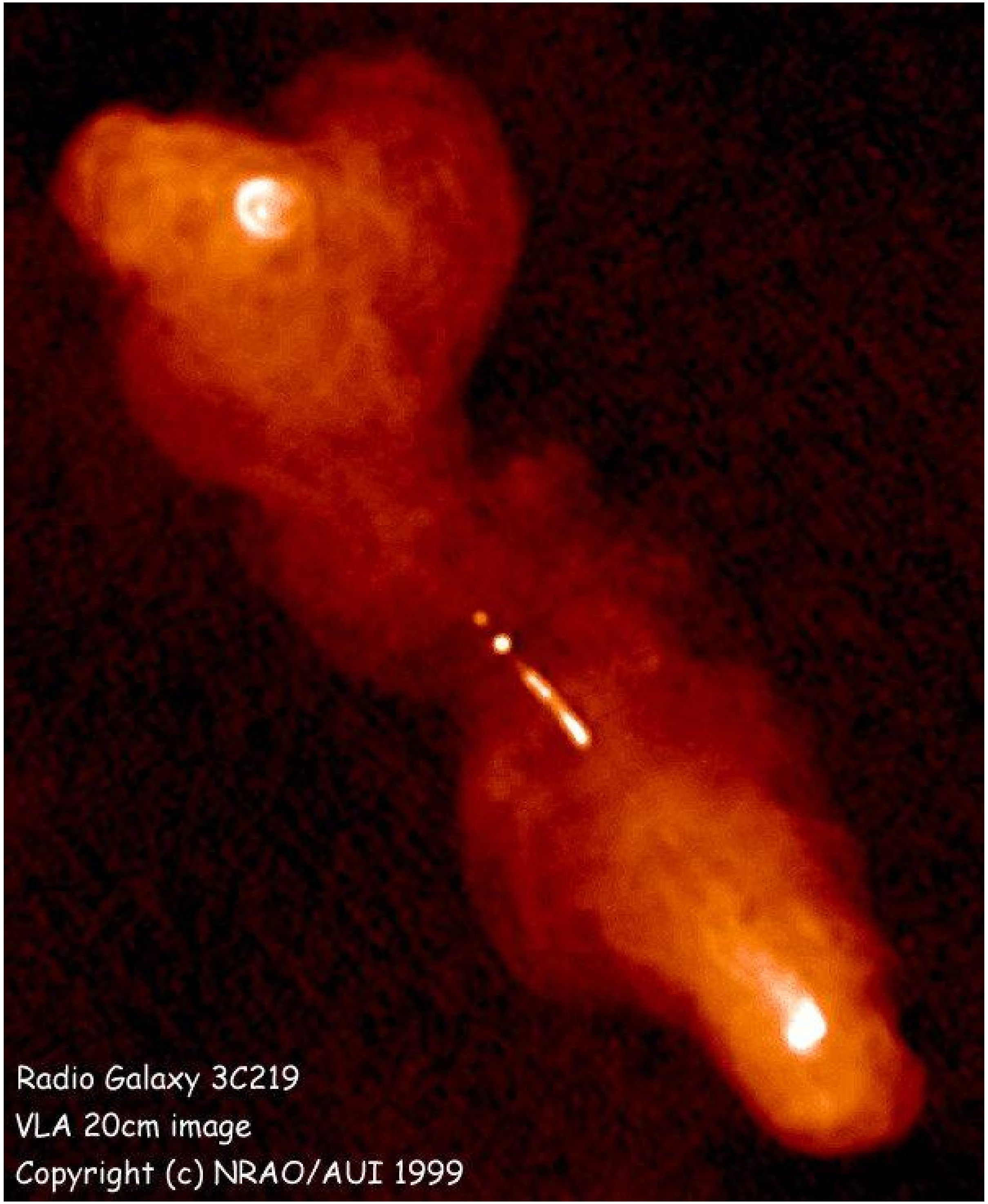


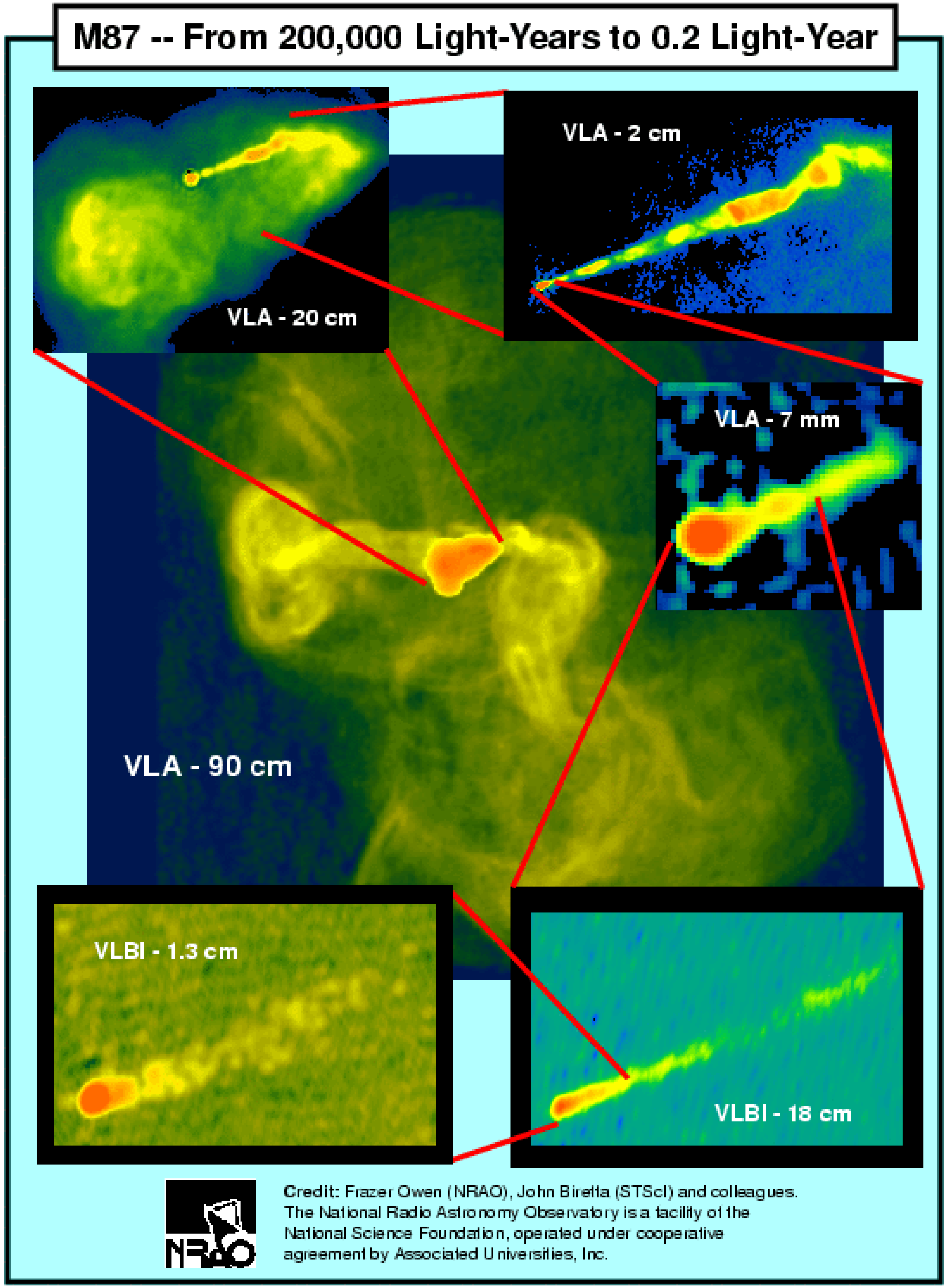
Cygnus A – X-ray image

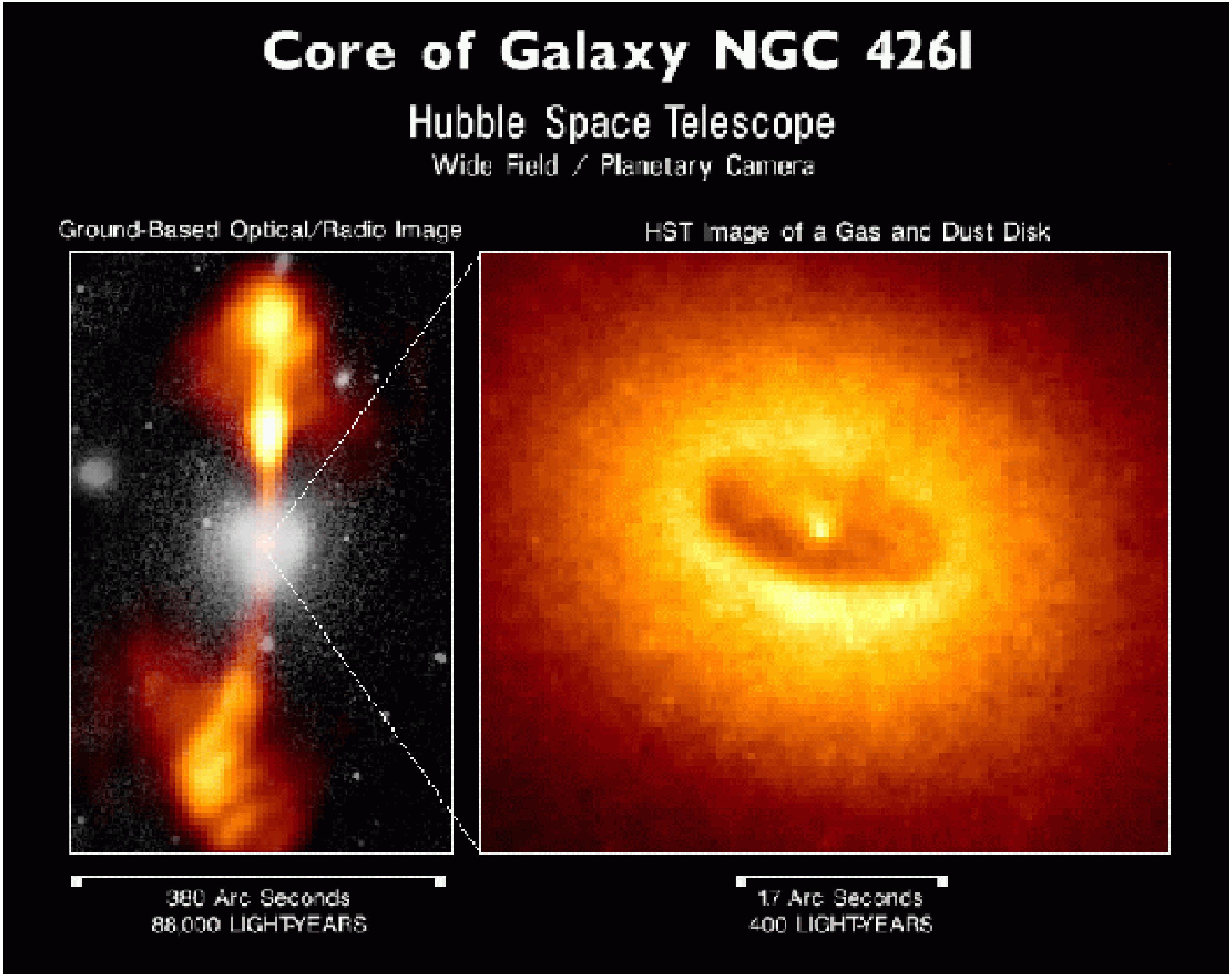
The image shows the distribution of hot intergalactic gas surrounding the radio source.




3C219
BLUE = V BAND
RED = 22 CM


Radio Galaxy 3C219
VLA 20cm image
Copyright (c) NRAO/AUI 1999

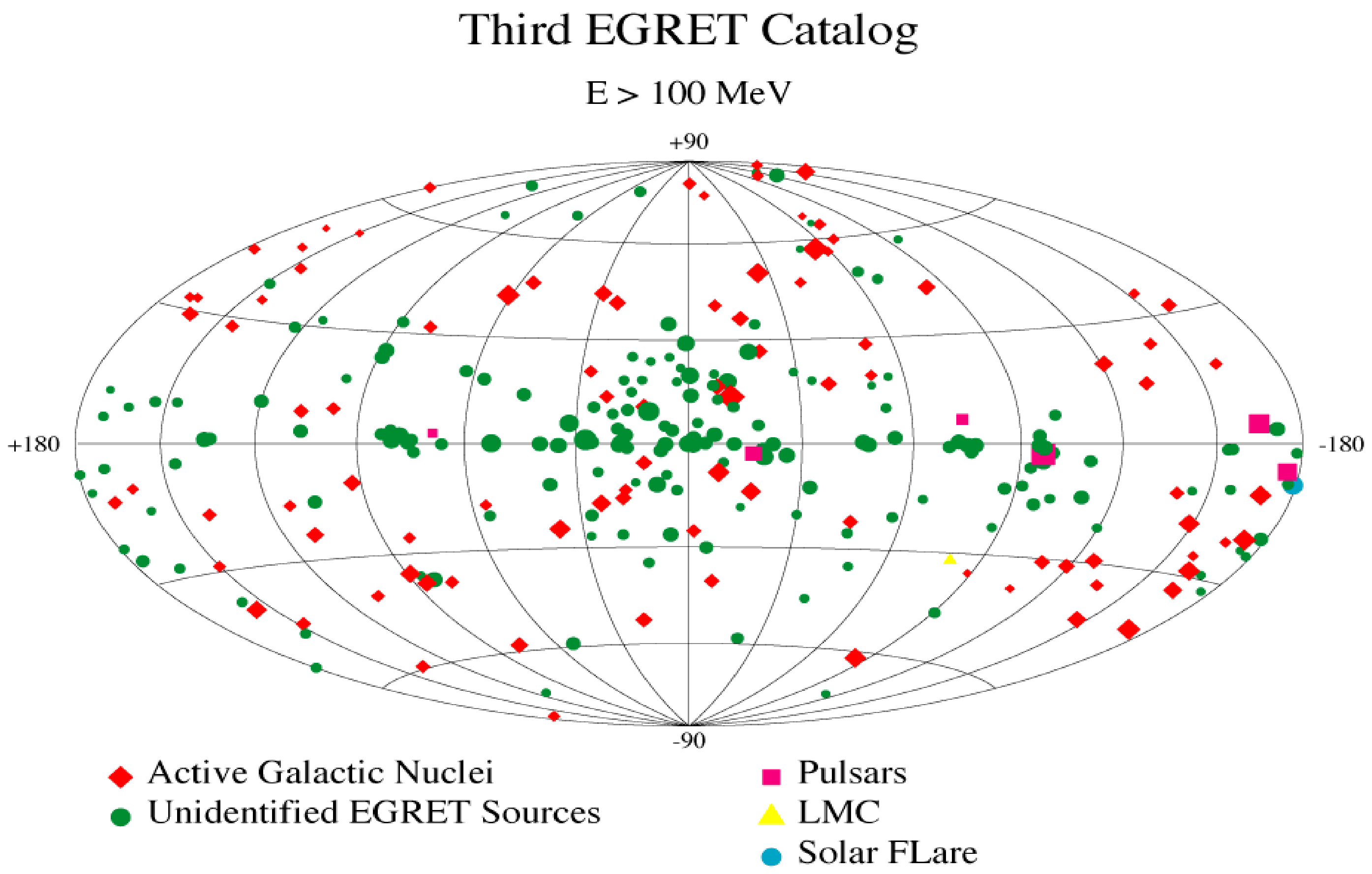

M87 -- From 200,000 Light-Years to 0.2 Light-Year
VLA - 20 cm
VLA - 2 cm
VLA - 7 mm
VLA - 90 cm
VLBI - 1.3 cm
VLBI - 18 cm
Credit: Frazer Owen (NRAO), John Biretta (STScI) and colleagues.
The National Radio Astronomy Observatory is a facility of the
National Science Foundation, operated under cooperative
agreement by Associated Universities, Inc.


Core of Galaxy NGC 4261
Hubble Space Telescope
Wide Field / Planetary Camera
Ground-Based Optical/Radio Image
HST Image of a Gas and Dust Disk
380 Arc Seconds
88,000 LIGHT-YEARS
17 Arc Seconds
400 LIGHT-YEARS

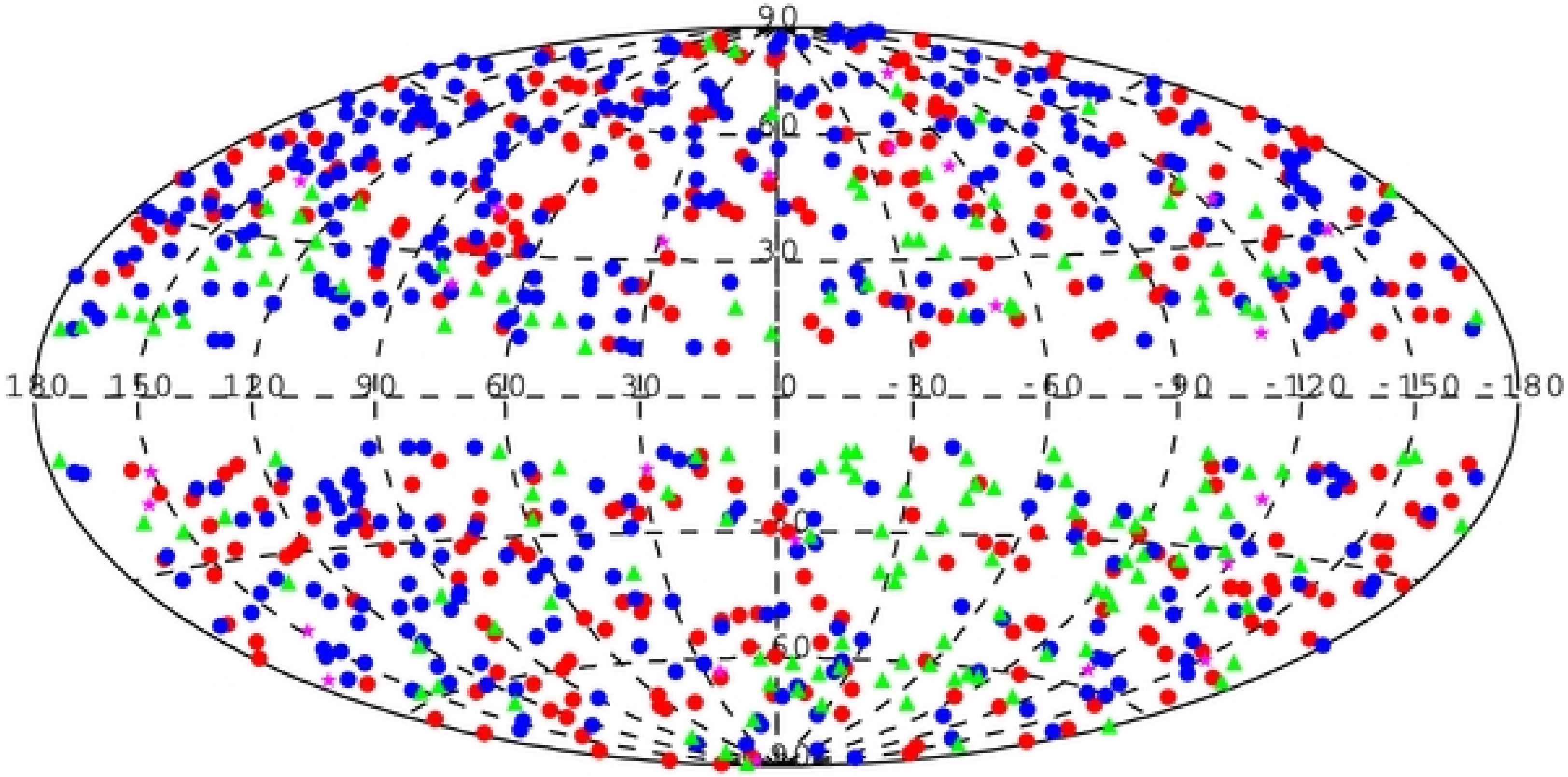
EGRET

- Before EGRET : 3C273
- After EGRET : 33 AGN a more than 5 ¾

  all Blazars



FERMI 2nd LAC catalogue



886 AGN with a clear GeV gamma-ray signal, identified with comprising 395 BL Lacertae objects (BL Lac objects), 310 flat-spectrum radio quasars (FSRQs), 157 candidate blazars of unknown type (i.e., with broadband blazar characteristics but with no optical spectral measurement yet), 8 misaligned AGNs, 4 narrow-line Seyfert 1 (NLS1s), 10 AGNs of other types, and 2 starburst galaxies

EGRET

- AGN basic results :

- The γ-ray energy flux is often dominent
- A naive estimate $(4\pi d^2 L) \approx 10^{47} \rightarrow 10^{49} ergs.s^{-1}$
- $0,03 \leq z \leq 2,28$
- $(30MeV - 30GeV)$spectrum $\sim$ a power law:
  $\frac{dN}{dE} \propto e^{-\alpha}$ with $1,3 \leq \alpha \leq 3,0$
- Many sources are strongly and rapidly variables on time scales $\sim$ 1 month.
- Many known AGN are not detected in γ-rays.

EGRET

- Variabilité:
  - La plupart des Blazars EGRET présentent une variabilité sur des échelles de temps comprises entre le jours et un mois.

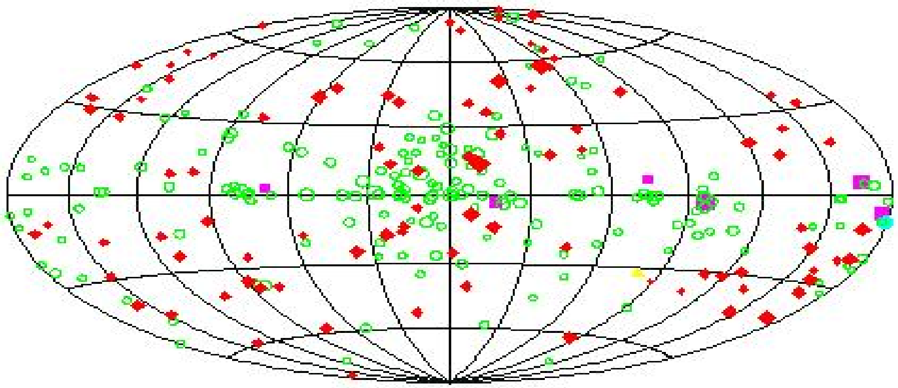
Source	T_{var} (jours)	T_{var} AGN (jours)
0208-512	8	4
0235+164	<90	<46
0446+112	15	...
0528+134	2	0,65
1253-55	2	1,3
1406-076	6	2,4
1633+382	2	0,7
2251+158	4	2,2

  - Les temps très courts contraignent la géométrie de la source :

    1 jour ⇒ Rayon de Scharzchild d'un trou noir de $10^{10} M_{\odot}$

Third EGRET Catalog

E > 100 MeV

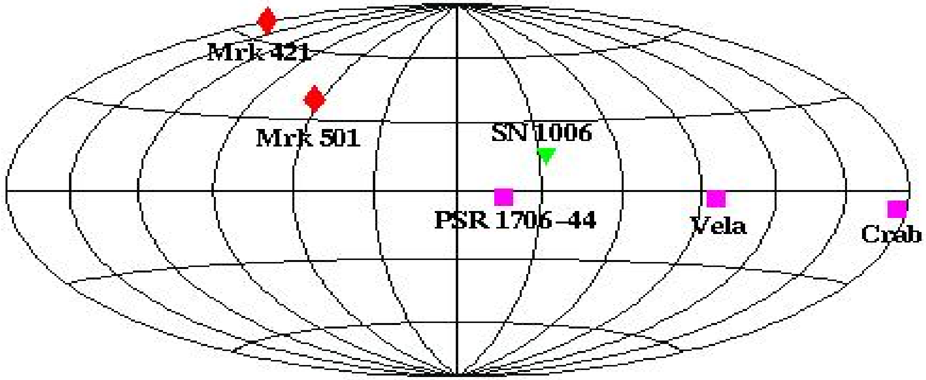


Non identified sources   7 Pulsars   5 Supernova remnants

Blazars (AGN)   LMC

Sources seen from the ground

E > 250 GeV

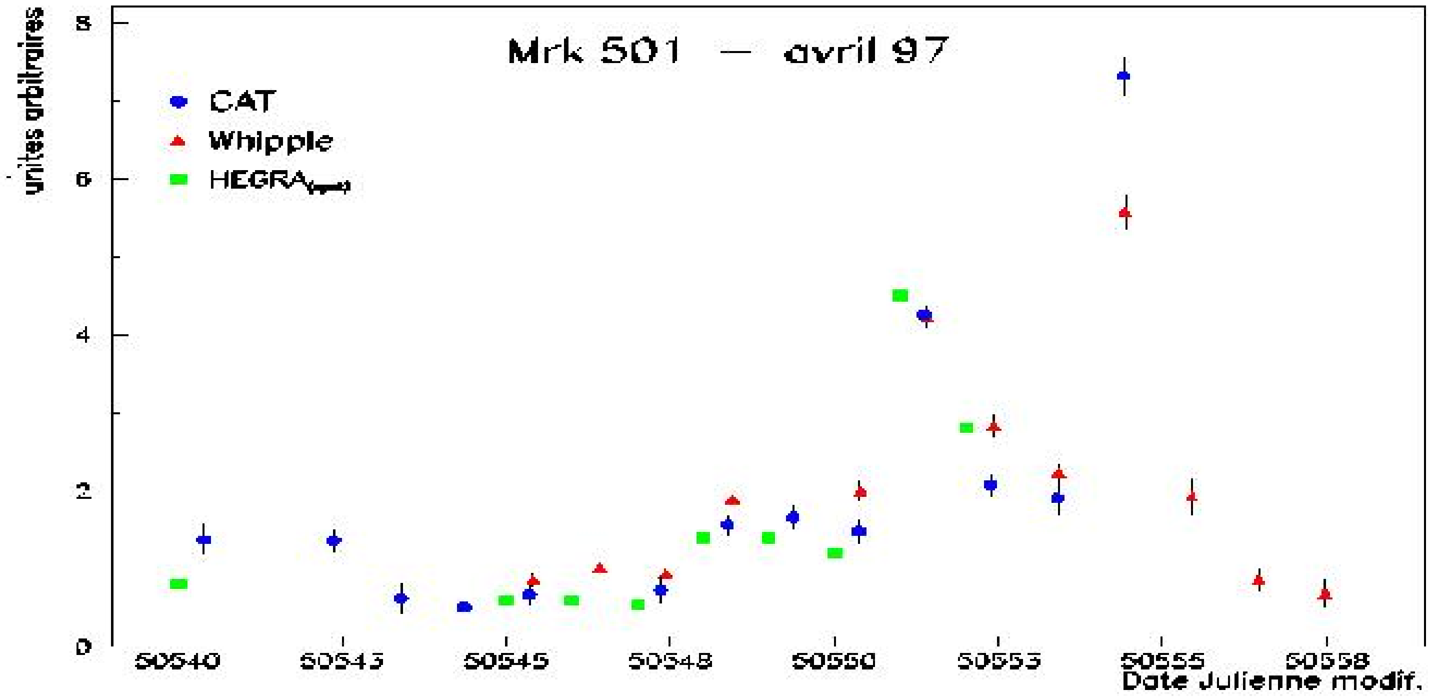


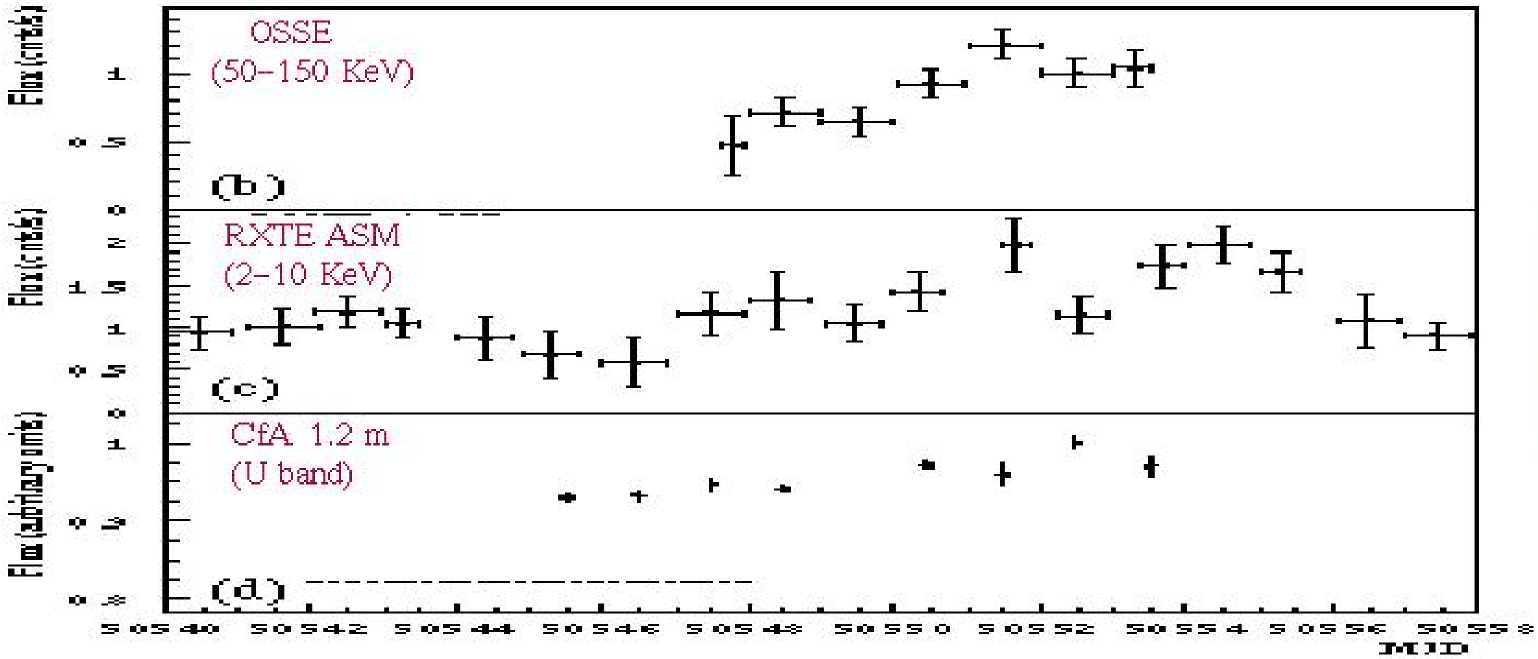
After COS-B, the great success of EGRET has been the rather unexpected discovery that ≈ 100 blazars are brilliant in the GeV energy range

with no turn-over at maximum energies

During about the same decade
the TeV observations broke through,
with no pre-notice.
The a posteriori suprise is the scarcity of sources

- the intermediate region, from 30 to 300 GeV, remains unexplored
- the (few) TeV blazars are weak EGRET sources
- the absorption ($\gamma\gamma \rightarrow e^+e^-$) obscures the far Universe for TeV γ's

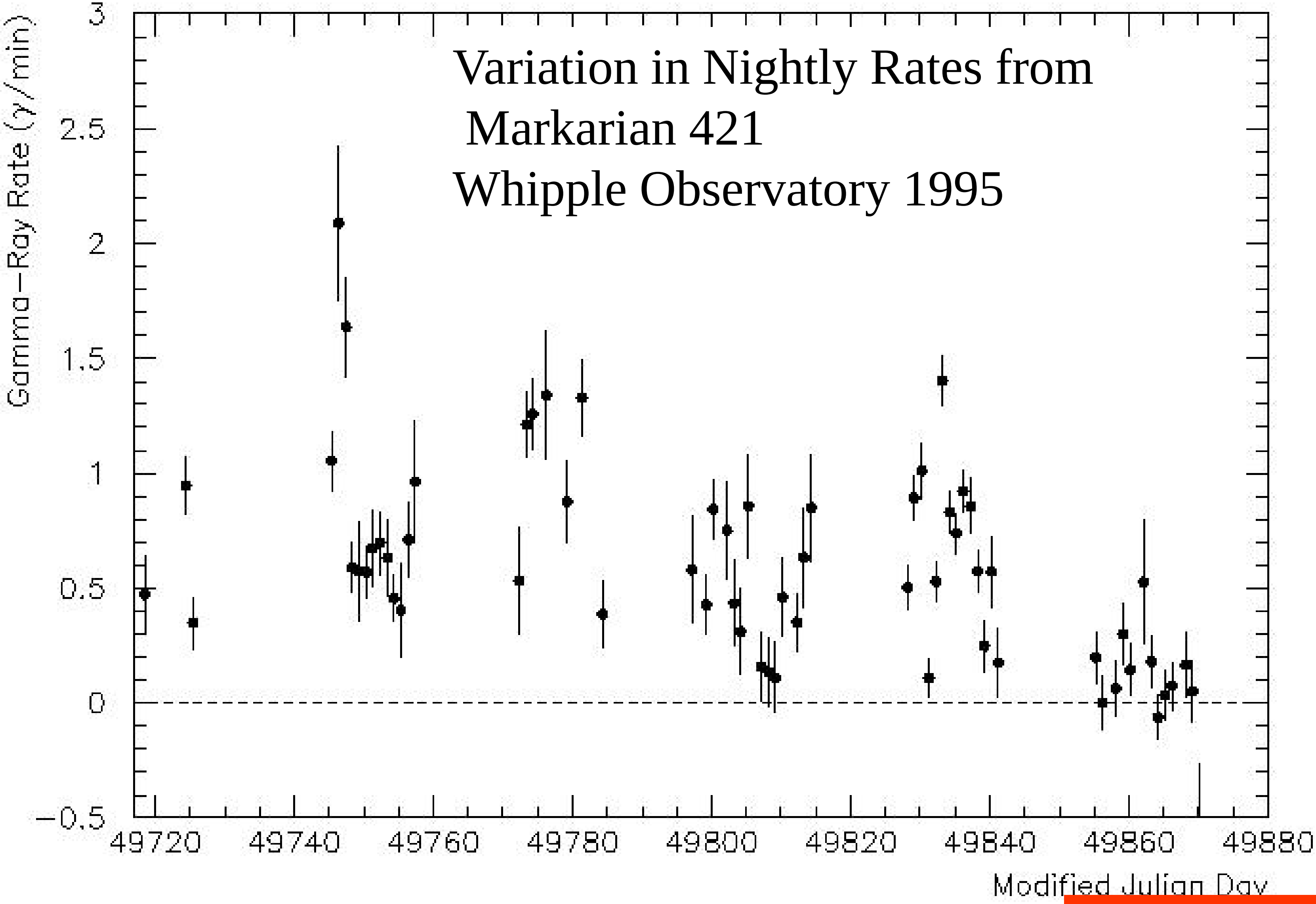




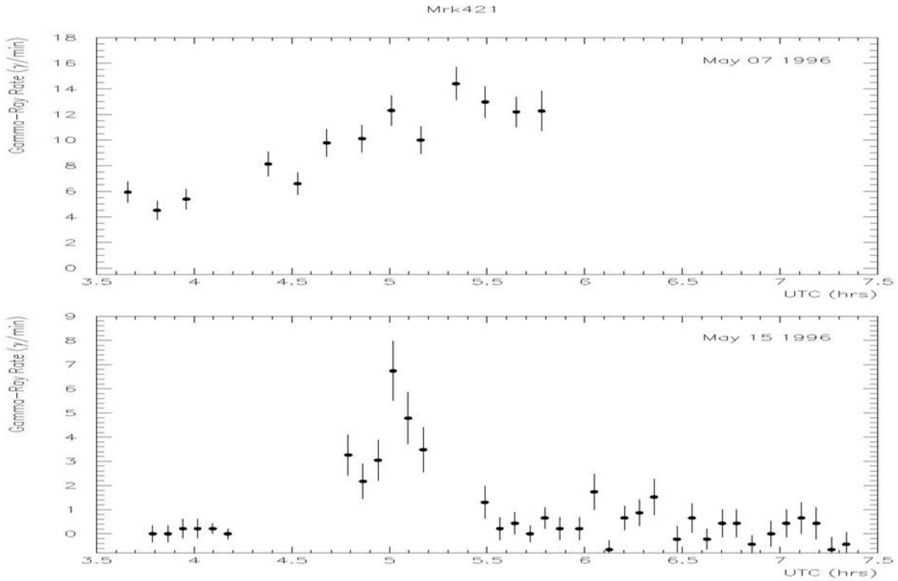
Mrk 501
flares of April 1997
showing variability
on a day scale

Even faster (<hr)
variations have
been evidenced by
the Whipple
on Mrk 421

Flux sensitivity
on the scale of
the hour is needed

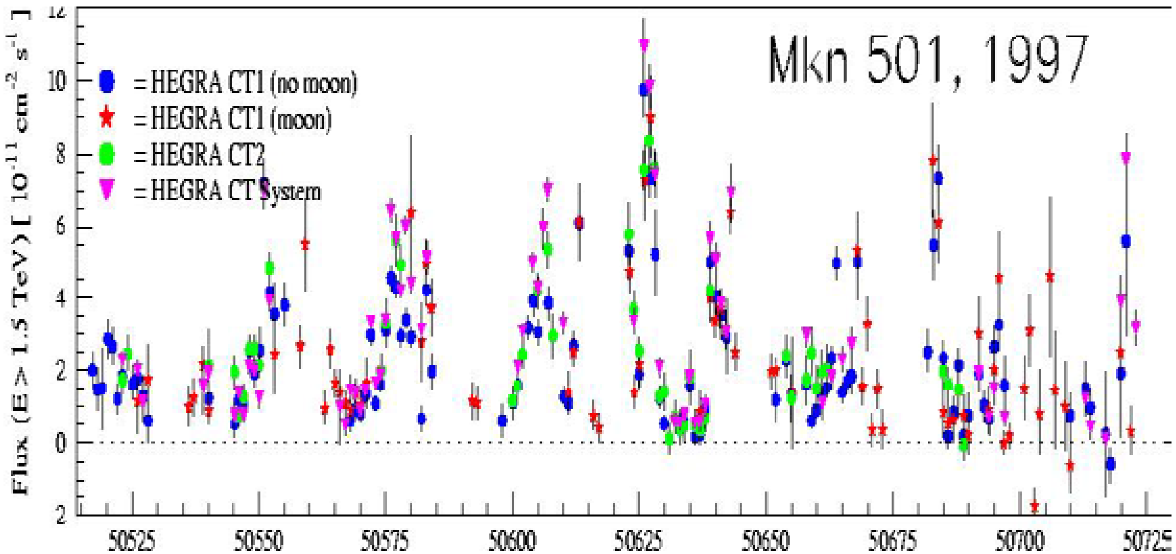


Buckley et al. 1996

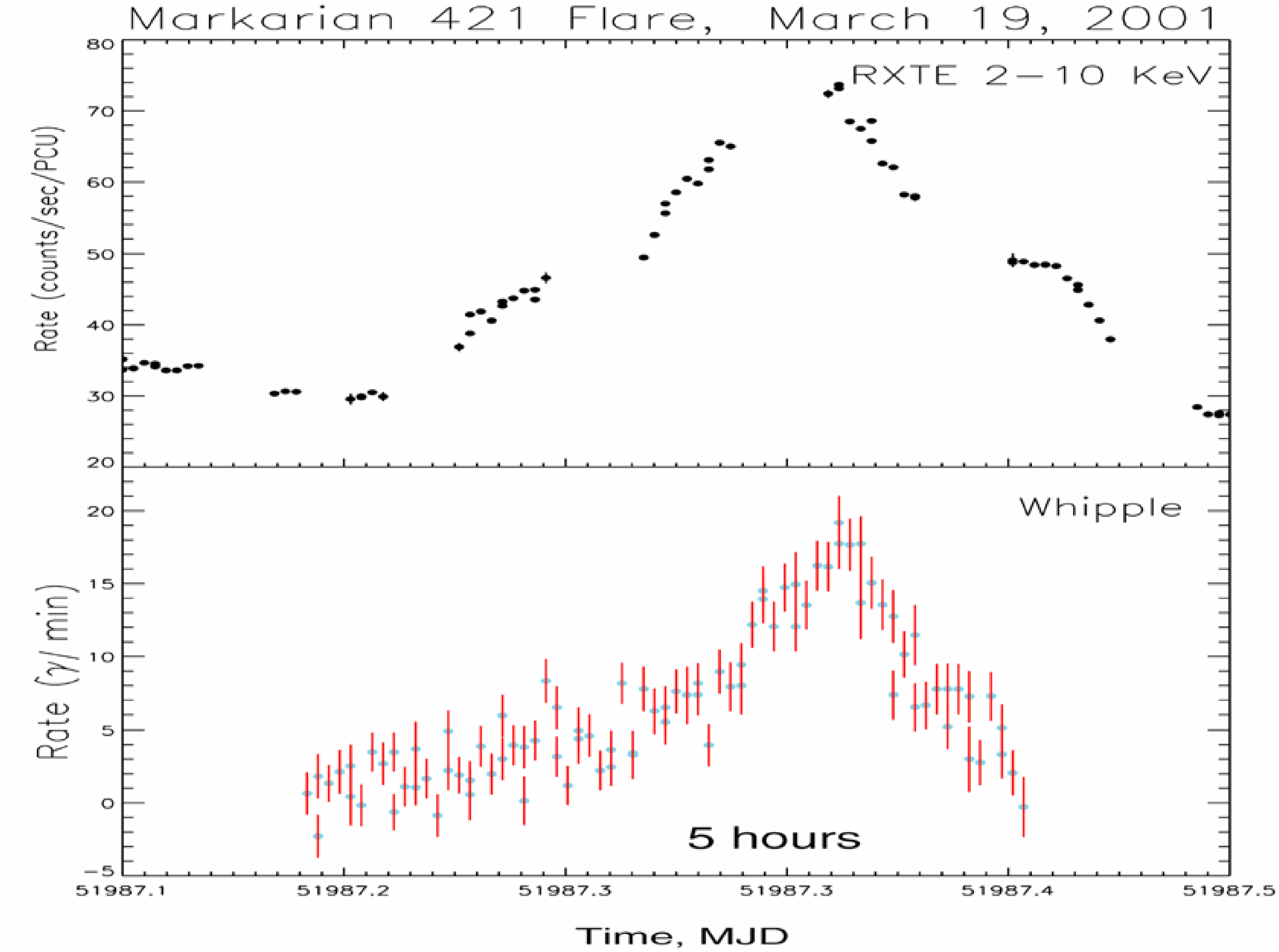


Markarian 421
Whipple Observatory
Hours - minutes

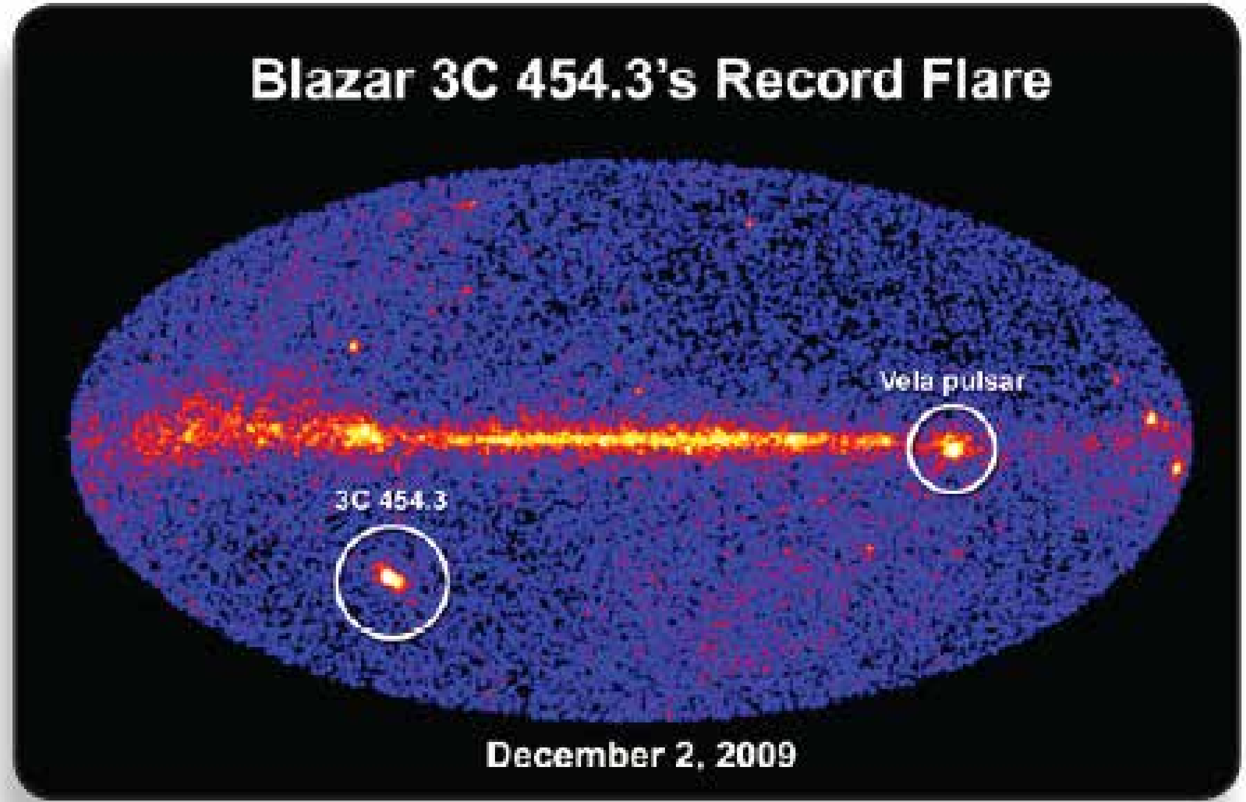
Markarian 501
HEGRA
Days-months



Time coincidence X and TeV g

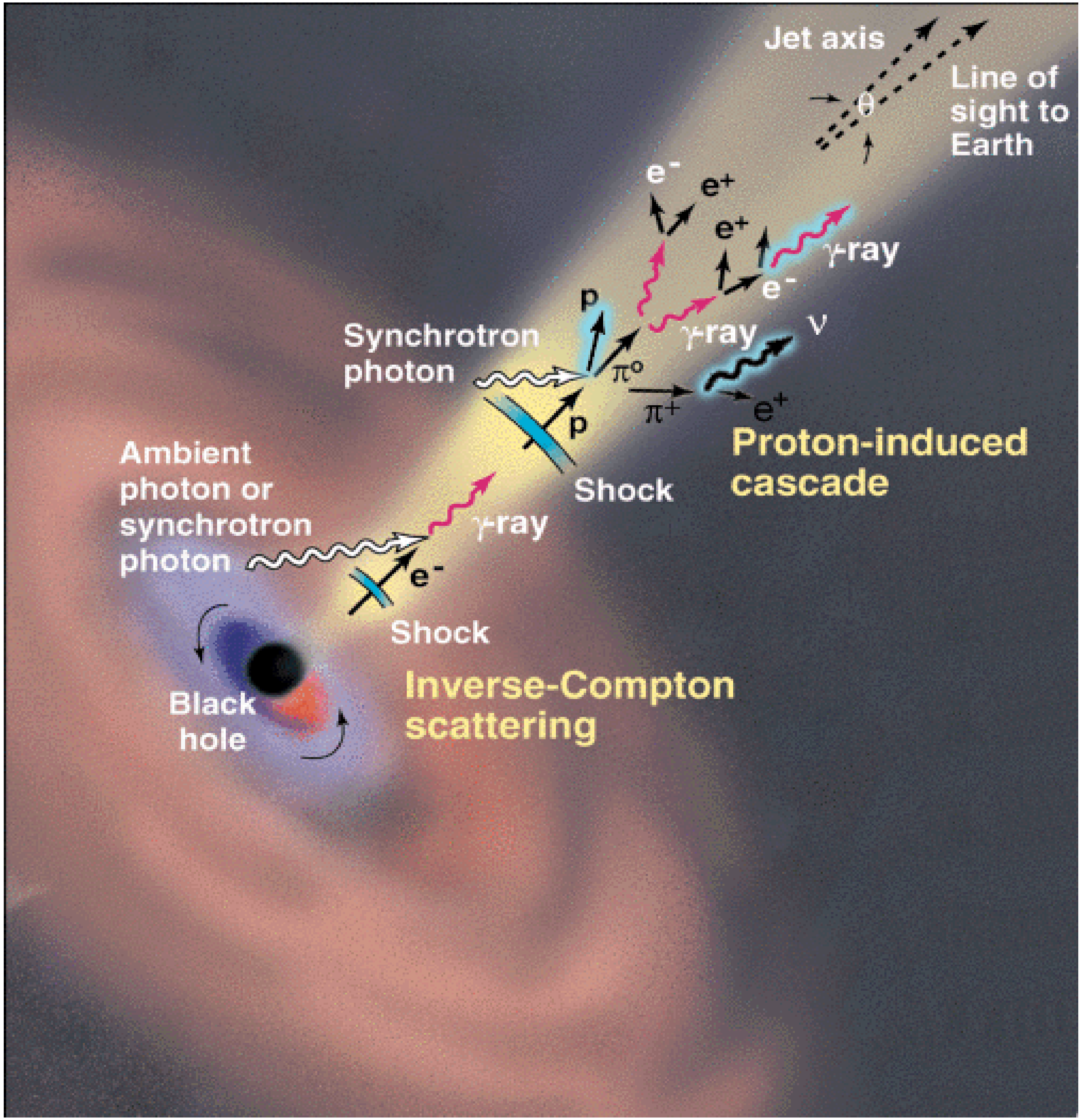


Flare de 3C454.3 : 2 x luminous than Vela although 10^6 x distant !



Unprecedented flares from the blazar 3C 454.3 in the constellation Pegasus now make it the brightest persistent gamma-ray source in the sky. That title usually goes to the Vela pulsar in our galaxy, which is millions of times closer (Credit: NASA/DOE/Fermi LAT Collaboration).

AGN Jet Emission Mechanisms



Cosmic Cannon:
Looking down the Barrel of the Cannon

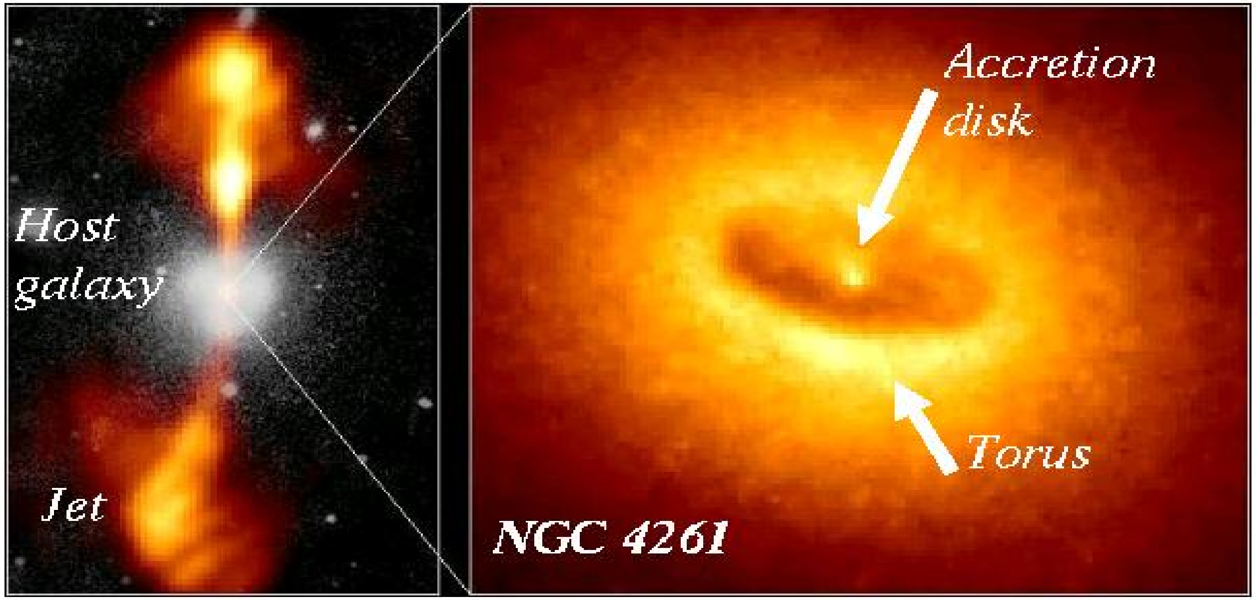
Electron Progenitors:
- **Synchrotron Self Compton**
- **External Compton**

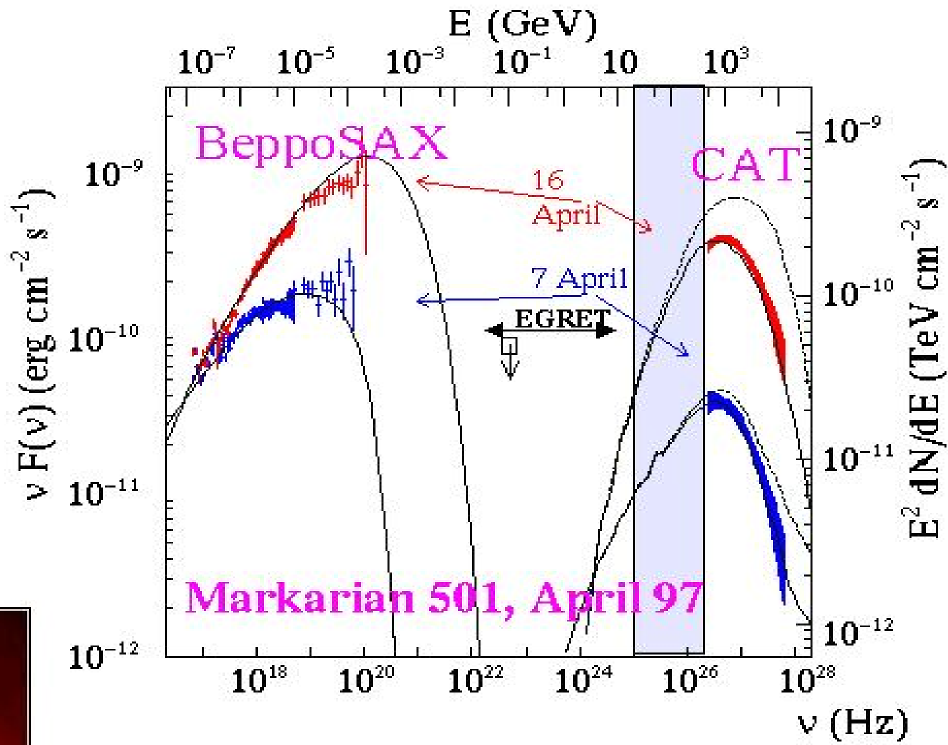
Proton Progenitors:
- **Proton Cascades**
- **Proton Synchrotron**

(Buckley, Science, 1998)

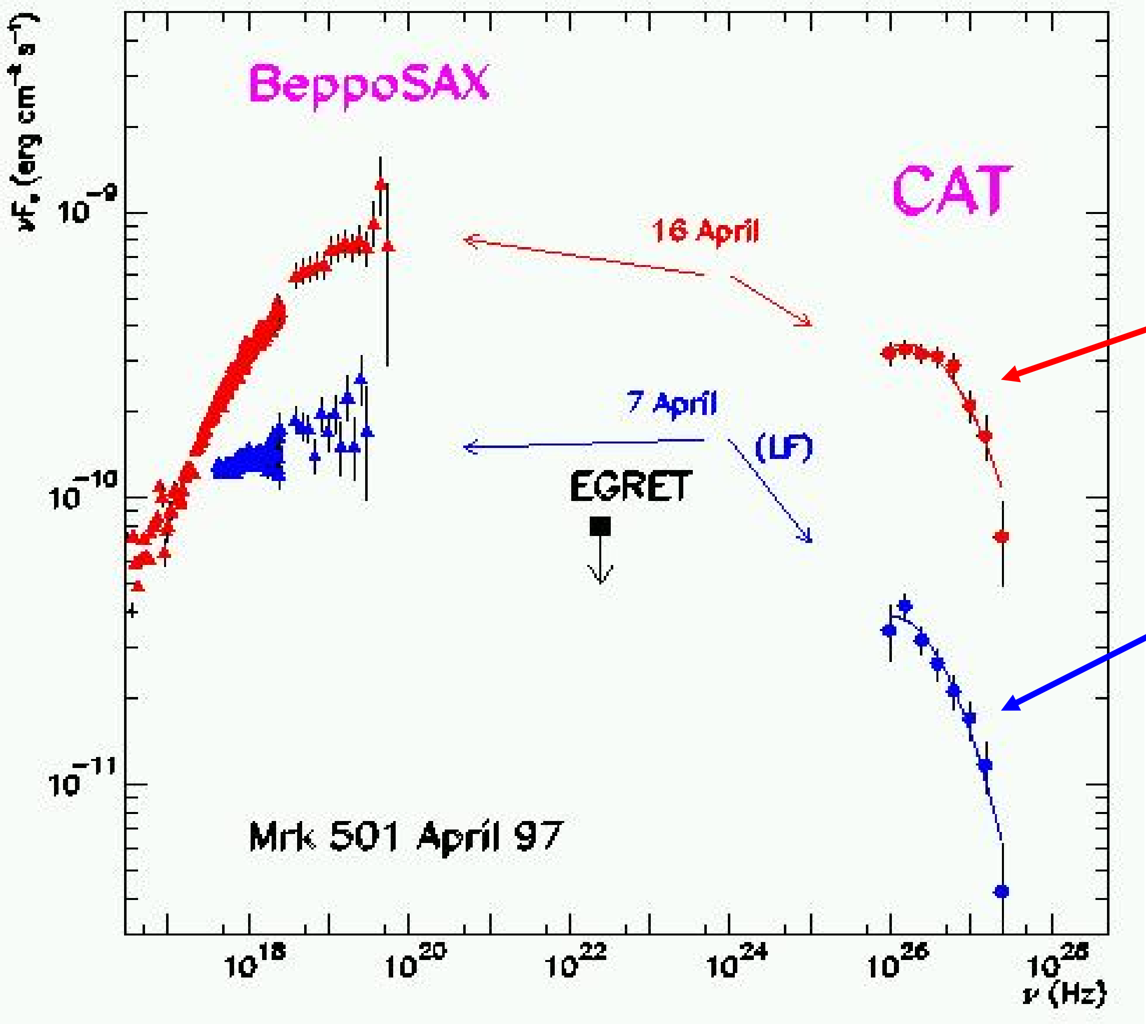
- AGN / Blazar : Markarian 501
  - Supermassive Black Hole ($M > 10^6 M_\odot$)
  - With an acretion disk (1 pc)
  - Surrounded by a torus of dust (100 pc)
- Radio loud:

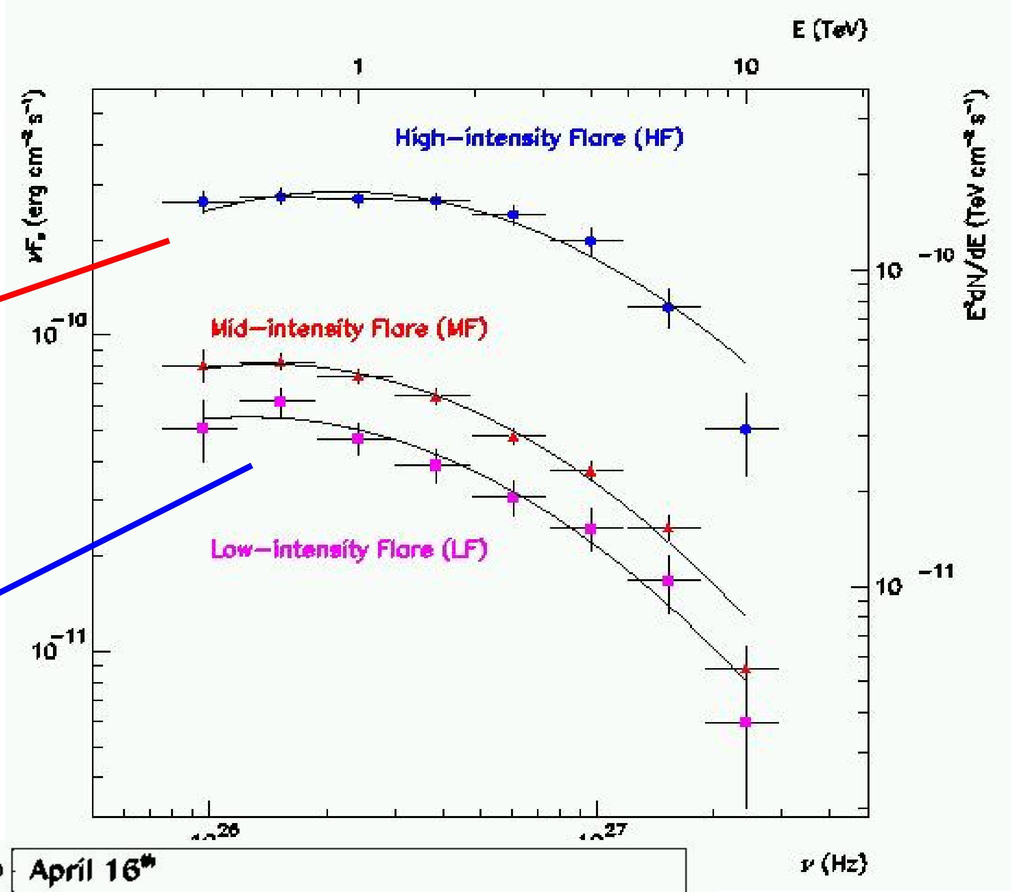
  Two jets (<100 kpc)

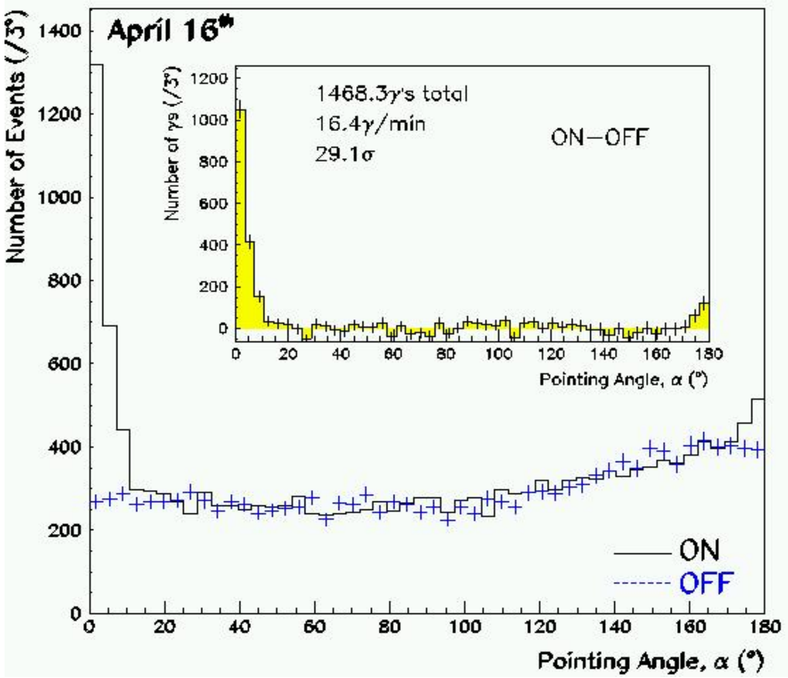




- *Blazars*
  - Near jet axis : $\theta \leq 1/\gamma$
    –> High energy emission
  - Strong variability on time scale of a *day* or even less





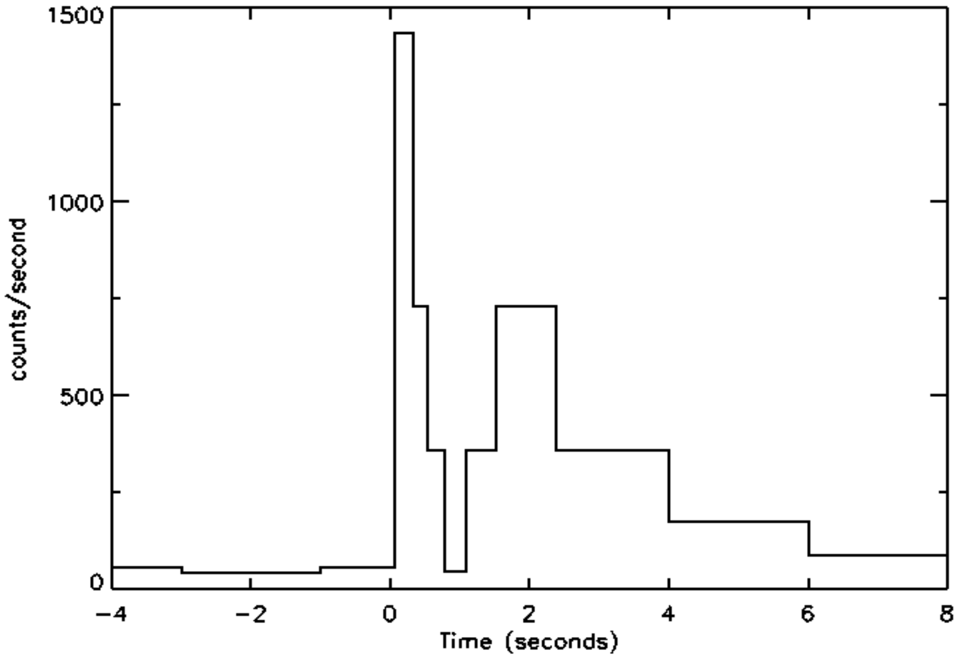


Markarian 501

Flare of April 1997

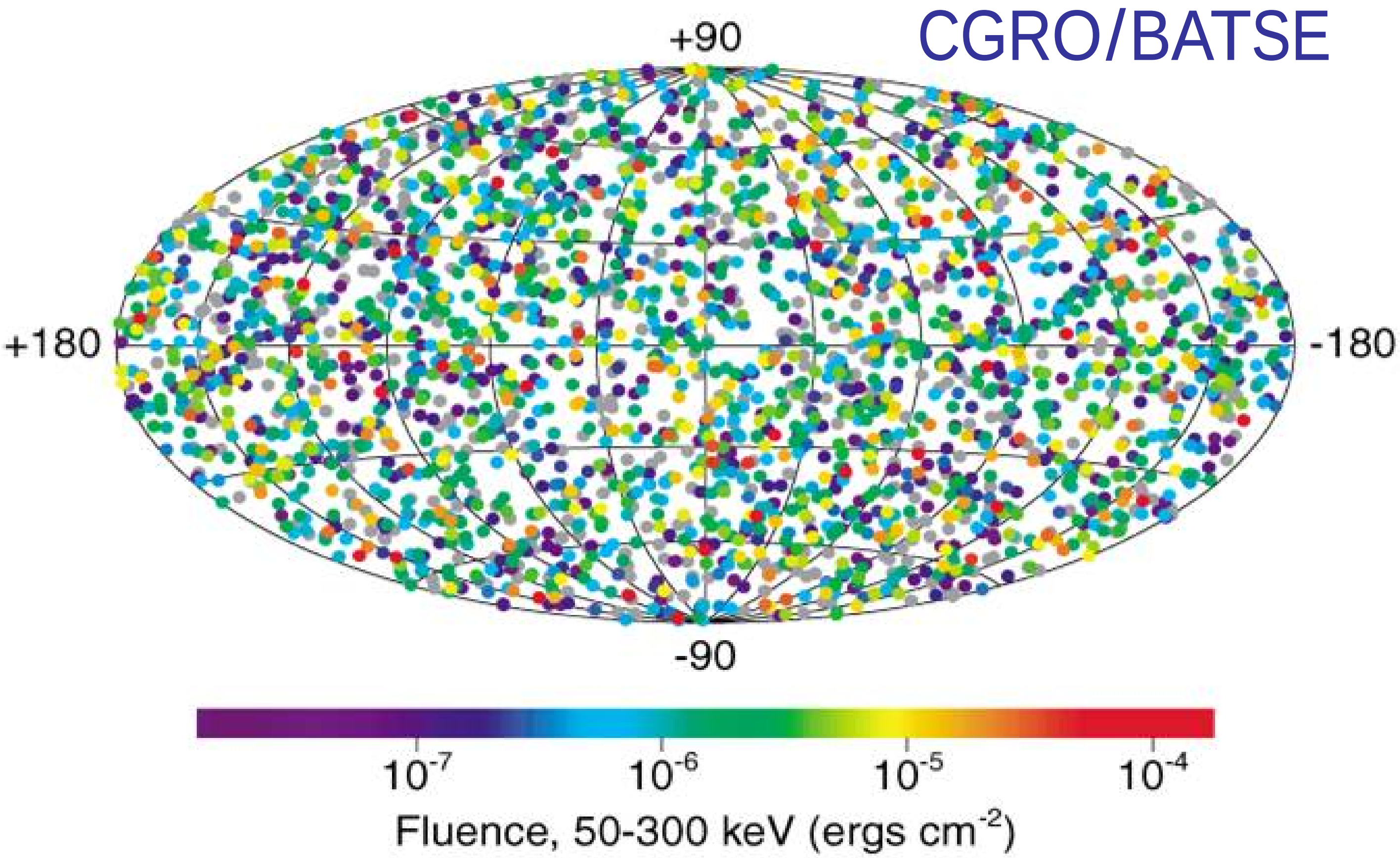
GAMMA RAY BURSTS

Vela Program (1969-1979)



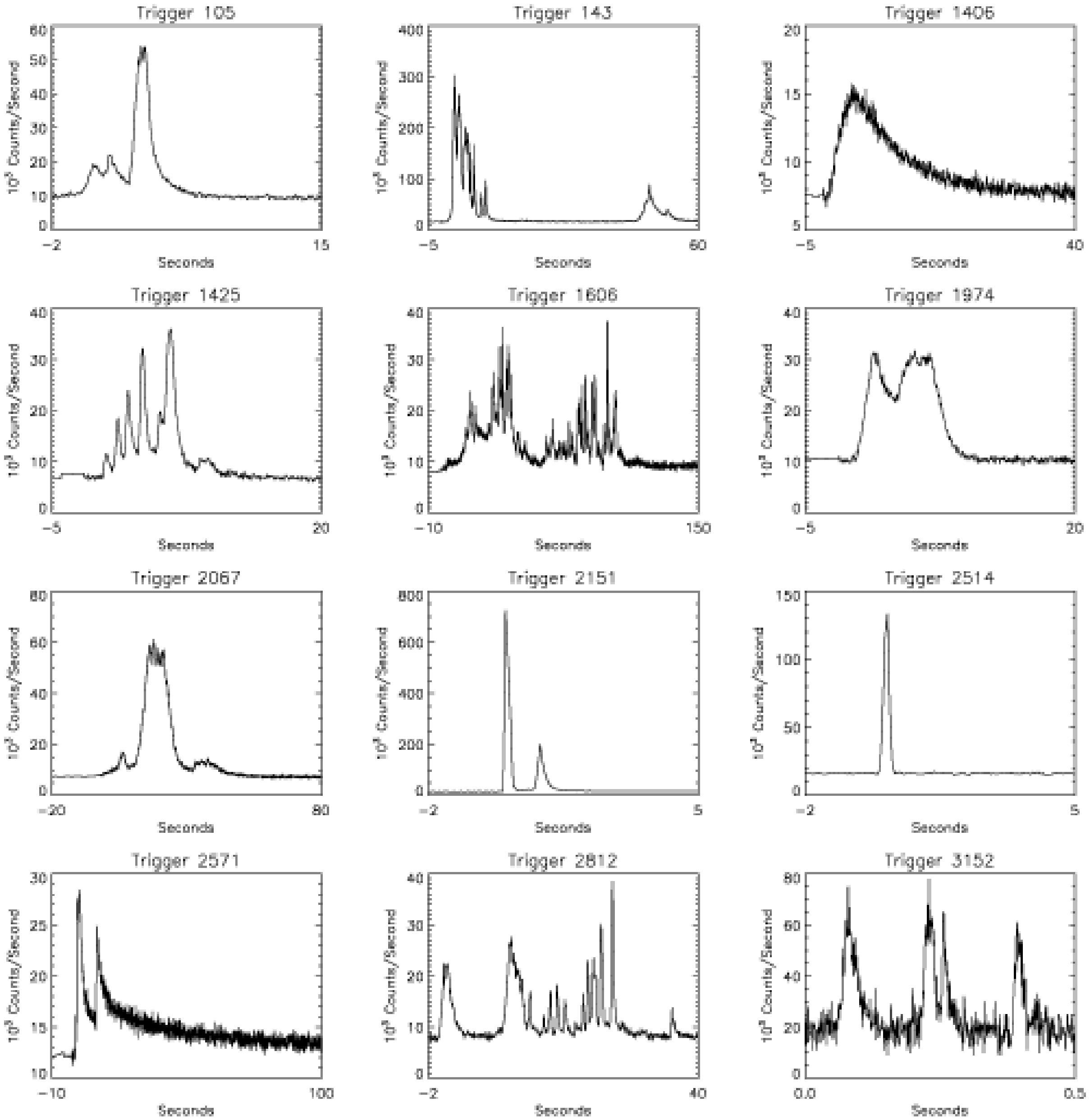
- Discovered in 1967 when spying thermonuclear bomb tests.
- A 30+ years old mystery unraveled in the 90ties !

Gamma-ray Burst Sky



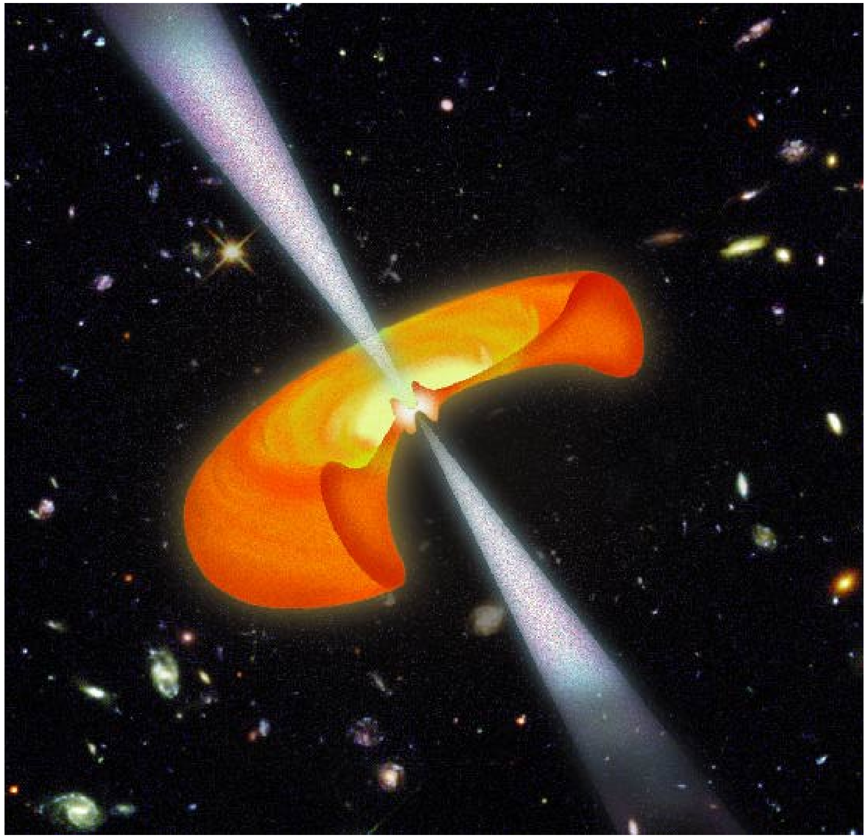
~ Once a day, anywhere in the universe !

GRB


Trigger 105
Trigger 143
Trigger 1406
Trigger 1425
Trigger 1606
Trigger 1974
Trigger 2067
Trigger 2151
Trigger 2514
Trigger 2571
Trigger 2812
Trigger 3152
Seconds
10^3 Counts/Second

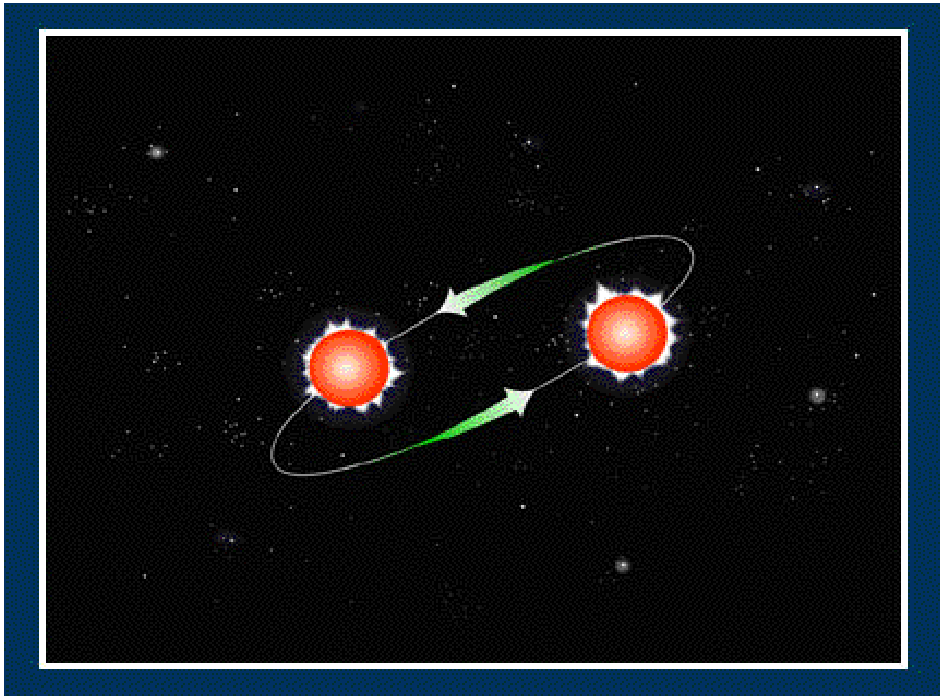
Hypernova

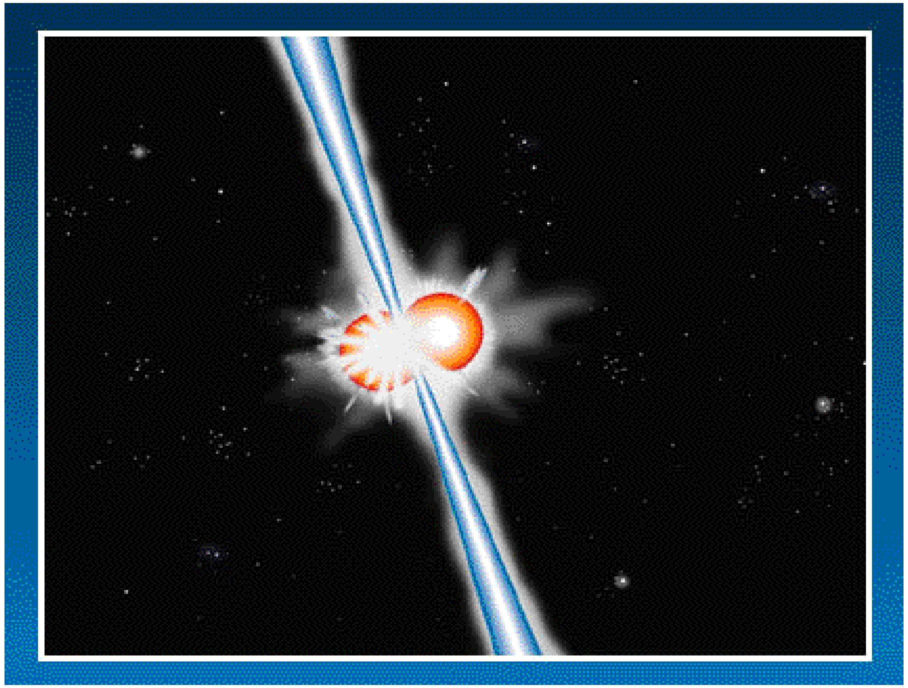
Death of a massive star ?



- A billion trillion times the power from the Sun

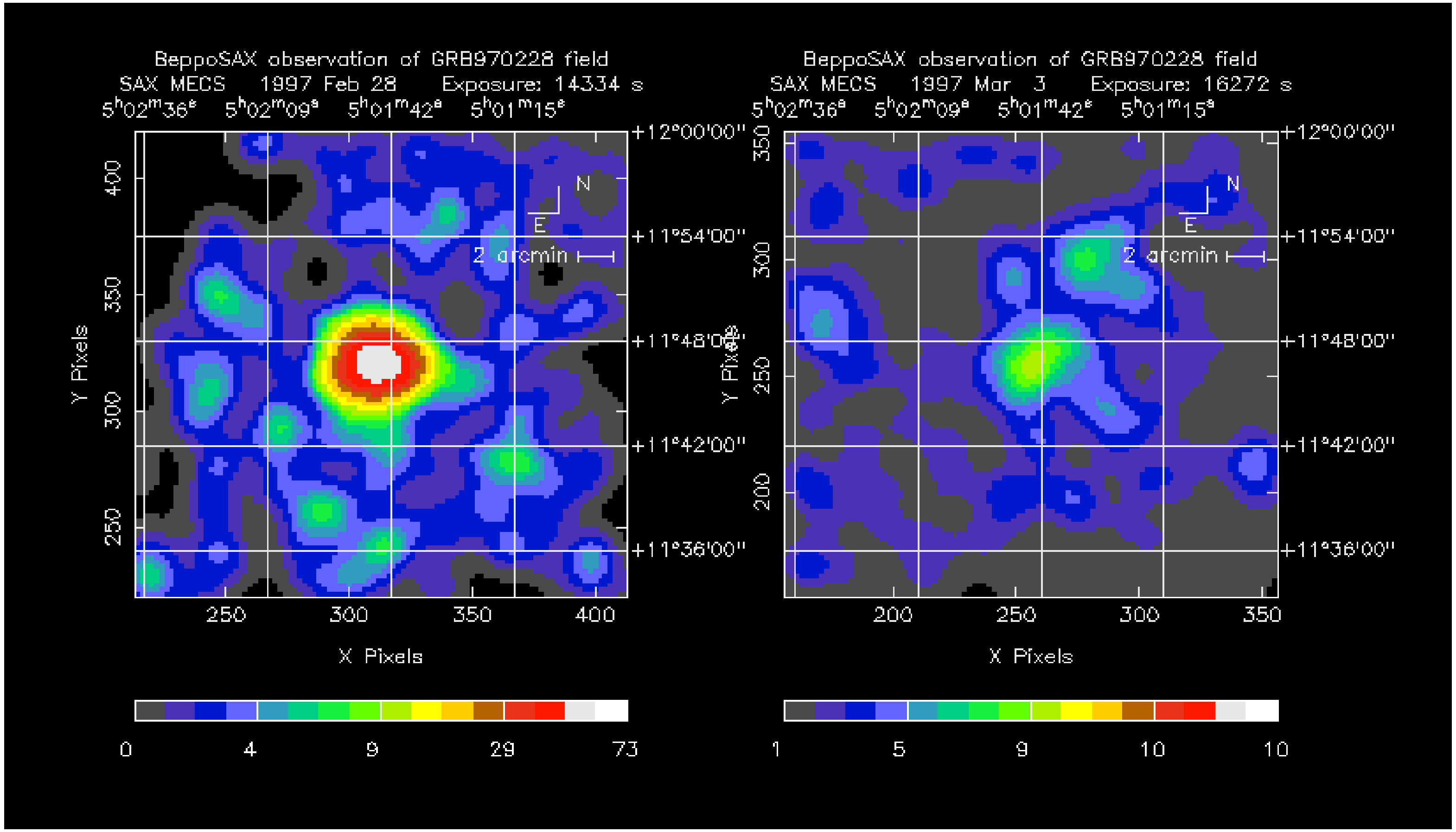
Catastrophic Mergers





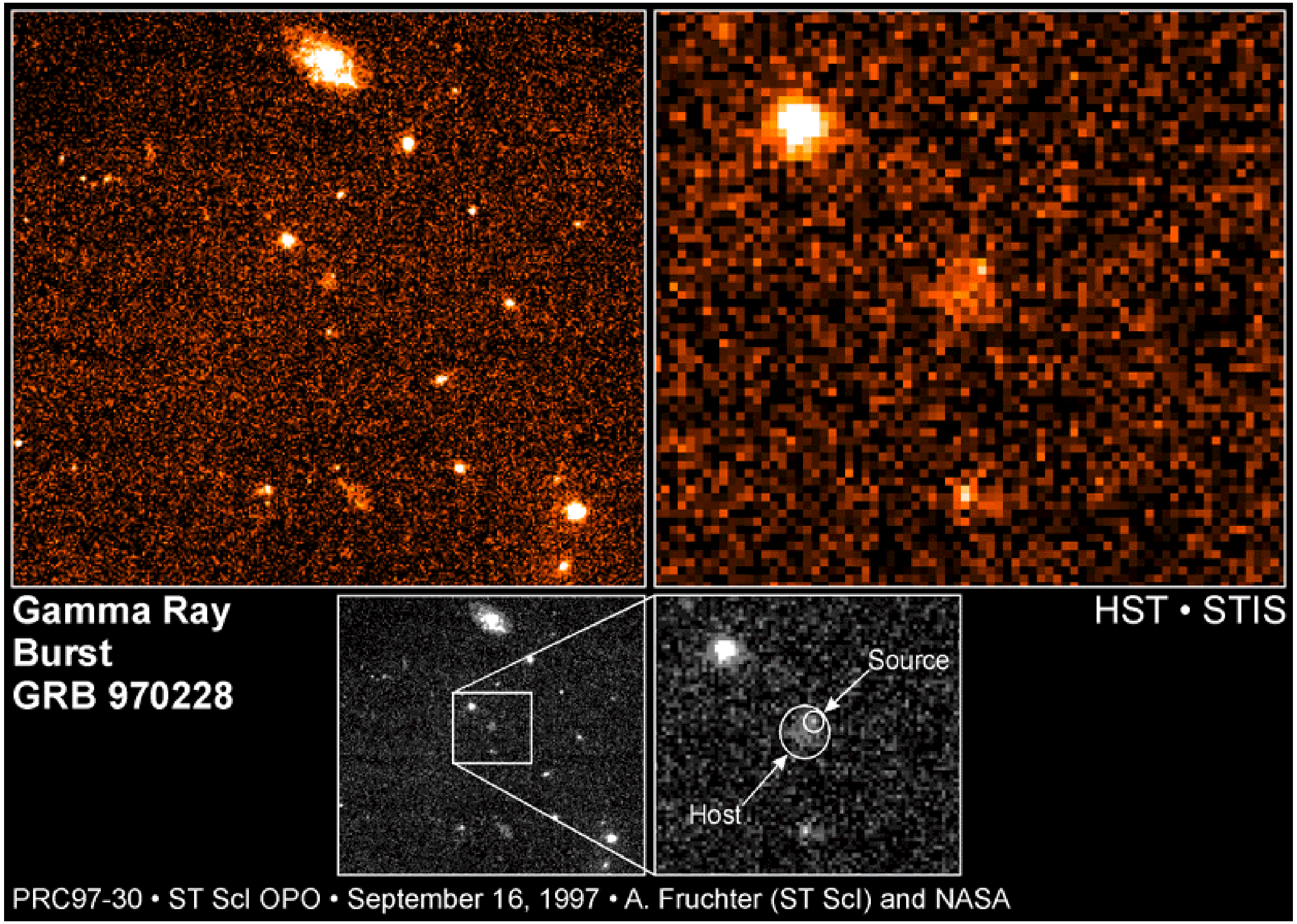
- Spiral death of 2 neutron stars or black holes

Afterglow



- Discovered in 1997 by BeppoSAX satellite

GRB afterglow en optique



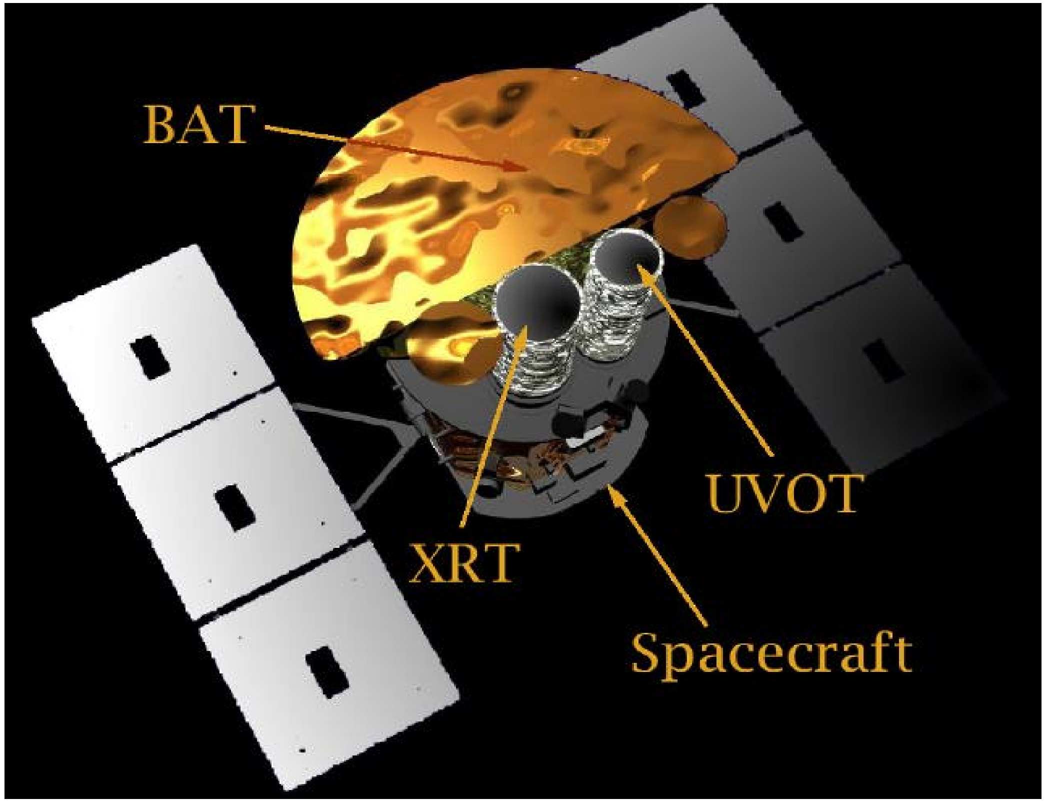
New Missions = Better Data

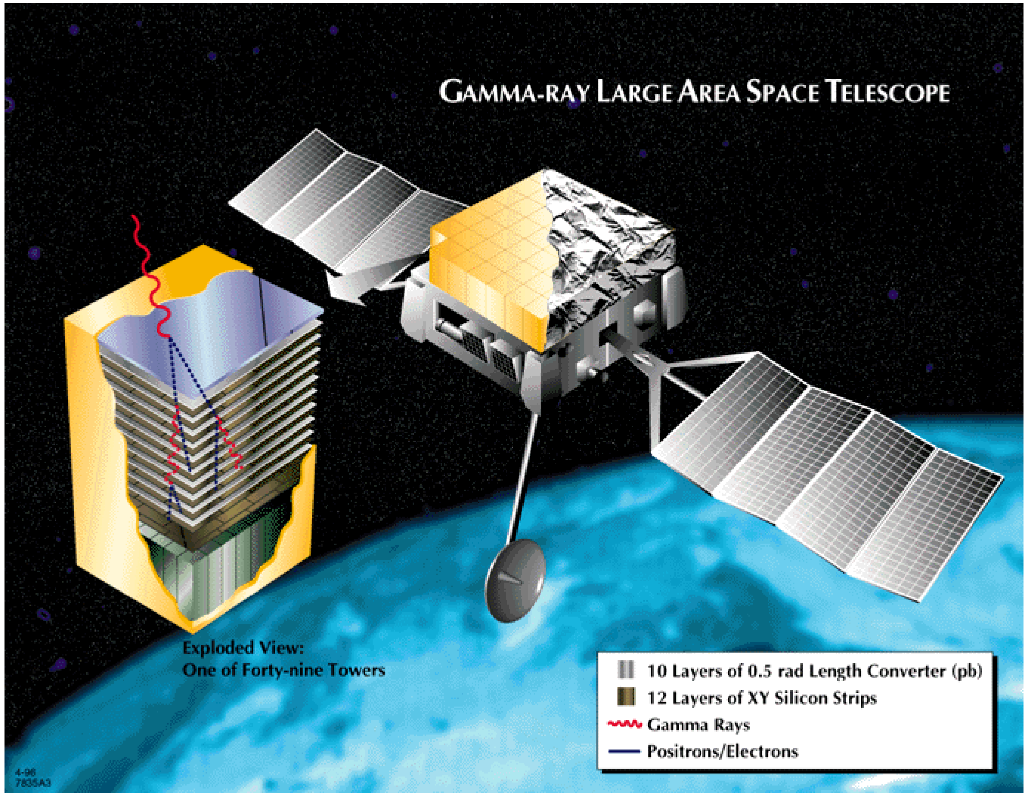
HETE II (launched 10/9/00)

INTEGRAL (launched 17/10/2002)

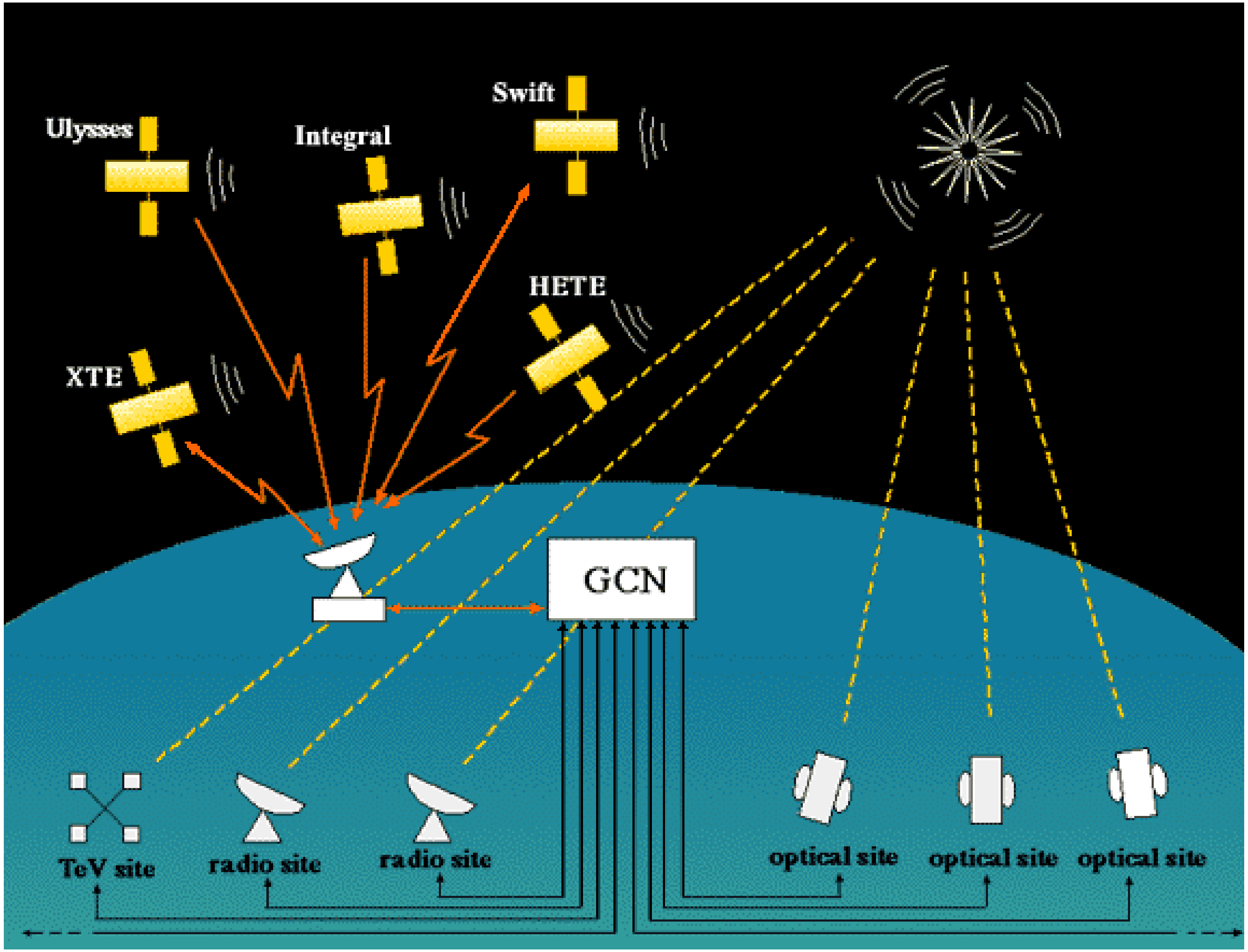
Swift (launched 20/11/2004)

FERMI (launched 16/5/2008)





The Gamma ray bursts Coordinates Network



GRB's

- Gamma-Ray Bursts : intense gamma-ray flux
  - 0.1 to 1 MeV, and up to 100 MeV
  - Emitted on a short time scale (~ 1 seconde !)
  - Observed ~once per day, with an isotropic distribution
  - Source at cosmological distances (most distant is z = 9.4 !!)

- °-ray luminosity: ~ 10^{52} erg/s (10 £ SNe !)

- Extreme variability in intensity and spectrum
  - Time scales from **10 ms** to **10 s**
  - Some very short **1 ms** variabilities ® internal shocks

- Clear bimodality suggesting the existence of two separate populations:
  - a "short" population with an average duration of about 0.3 seconds
  - a "long" population with an average duration of about 30 seconds.

UHE SOURCES

General limits on models

Shock waves

- Acceleration site confinment: $r_g = E/ZeBc < L$
  - Depends in fact on $V_{shock} = bc$:

$$E_{max} \approx b ´ \; Ze ´ \; Bc ´ \; L$$

Unipolar induction

- Accelerate in one step ($\mathbf{E}$ field, $f_{Lorentz} \approx V_{rot} ´ \; B ´ \; R$)

  ® No confinement necessary

Top-Down Models

- No acceleration at all !
  - Decay products of exotic physics states, supermassive particles at E_{GUT}, Topological Defects...

Bottom-Up

ZeVatrons

Astrophysical Accelerators reaching ZeV

Acceleration = Fermi-like diffuse acceleration.

Frist challenge, E_{max} : Reach ≥ 10^{20} eV

(if 1 TeV is hard, guess for 10^{9} TeV !!)

Second challenge, Propagation : B_{igm}

(determine the spectrum and the arrival)

Hillas criterion:

Magnetic confinement in the shock zone i.e.

1 R_{gyr} < accelerator size

Not many candidates survive!

Neutron stars (pulsars)
AGNs
Radio lobes
Clusters
Colliding Galaxies/Clusters
Gamma Ray Bursts

Hillas diagram

F.Montanet Cours Astroparticules M2R PSA & AMD Grenoble 2009-2010

Standard estimates for E_{max} :

- Confinement :

  $r_g = E/(ZeB) < R \Rightarrow E < ZeBR$
- Unipolar inductor (pulsar)

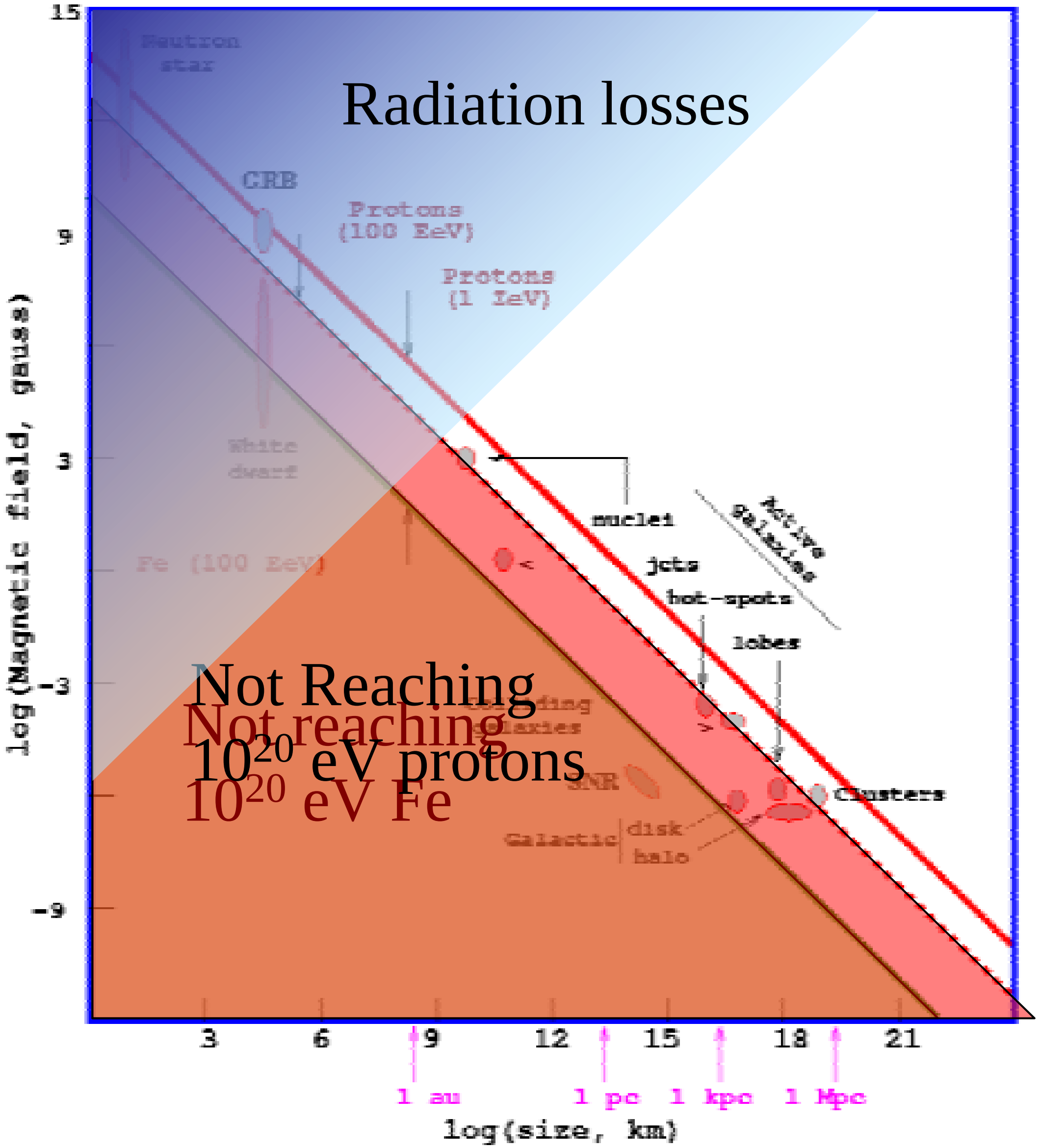
  $E < ZeBR(\Omega R/c) \approx \beta_s ZeBR$
- Diffusive acceleration by non relativistic shocks:

  $\tau_{acc} \approx 10\kappa/u_s^2 < R/u_s$ avec $\kappa > r_g c$

  $\Rightarrow E < \beta_s ZeBR$
- Diffusive acceleration by relativistic shocks:

  $E < \Gamma_s ZeBR$
- General Hillas dondition:

  $E < 0.9\beta\Gamma ZeB_{Gauss}R_{pc}\ ZeV$



$E_{max} \sim ZBL$ (Fermi)

$E_{max} \sim ZBL\Gamma$ (Ultra-relativistic shocks-GRB

Relativistic shocks

- Acceleration / G^2 works fine for a couple of cycle
- After that it fails for mere kinematical reasons
- But this is still very efficient (>> standard shocks)
- Confinement is easier

A weak deflexion is enough : $\delta\theta \approx 1/\Gamma_s \Rightarrow r_g < R_s/\Gamma_s$

$\Rightarrow E_{max} \approx \Gamma_s \times$ larger

$\Rightarrow$ one can reach the limits induced by energy losses

BOTTOM -UP

Galactic pulsars

Extragalactiques radio galaxy lobes

Gamma Ray Bursts

Protons, Iron, Nuclei?

Spectral index

Explaining isotropy is not trivial

Angular coincidences to be confirmed…

Top-Down

Topological defects, superheavy relics with M ~ GUT scale that is ~10^{16}GeV

- **Energy À 10^{20} eV easy!**
  (QCD fragmentation spectrum QCD with M~10^{24} eV‼)

- **Explaining the flux is not trivial ‼**
  (natural density scale is ~ H_0^{-1})

- **Composition of UHECR is the clue (photons + neutrinos) ‼**

Low energy gamma-rays constraint

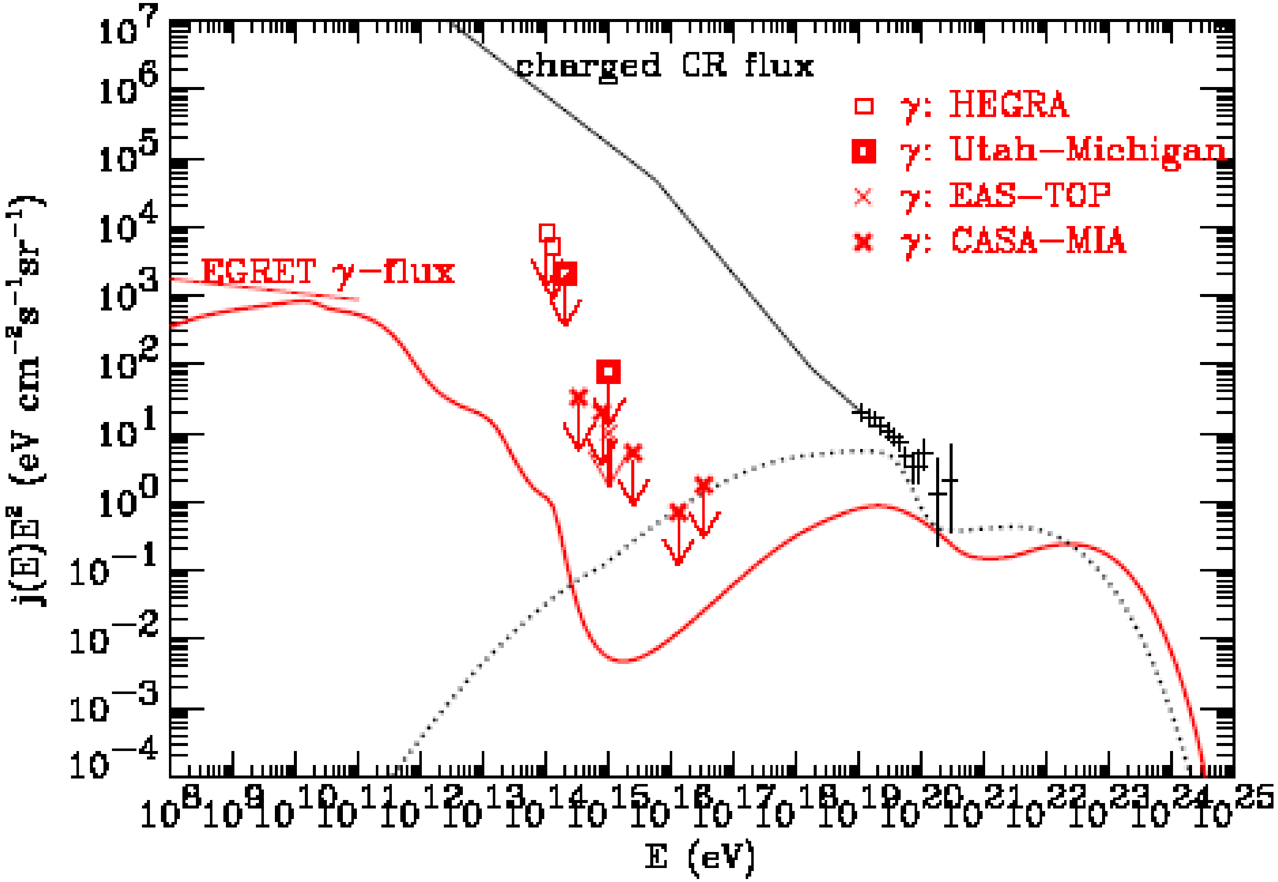


Figure 28: Predictions for the differential fluxes of γ−rays (solid line) and protons and neutrons (dotted line) in a TD model characterized by $p = 1$, $m_X = 10^{16}$ GeV, and the decay mode $X \to q + q$, assuming the supersymmetric modification of the fragmentation function, Eq. (57), with a fraction of about 10% nucleons. The calculation used the code described in Ref. [206] and assumed the strongest URB version shown in Fig. 10 and an EGMF $\ll 10^{-11}$ G. 1 sigma error bars are the combined data from the Haverah Park [3], the Fly's Eye [7], and the AGASA [8] experiments above 10^{19} eV. Also shown are piecewise power law fits to the observed charged CR flux (thick solid line) and the EGRET measurement of the diffuse γ−ray flux between 30 MeV and 100 GeV [185] (solid line on left margin). Points with arrows represent upper limits on the γ−ray flux from the HEGRA [257], the Utah-Michigan [510], the EAS-TOP [511], and the CASA-MIA [258] experiments, as indicated.

Top-Down Signatures

Composition:

Flux de Photons,Neutrinos À Protons

The current (AUGER) limits on UHE neutrino and photon flux already kill most Top-Down models !!

Spectrum:

QCD-like fragmentation spectrum quite "hard"

Cosmography:

Halo distribution!! (SHRs & TDs locales)

or ~ Homogeneous

and even more exotic stuff…

Strongly interacting neutrinos

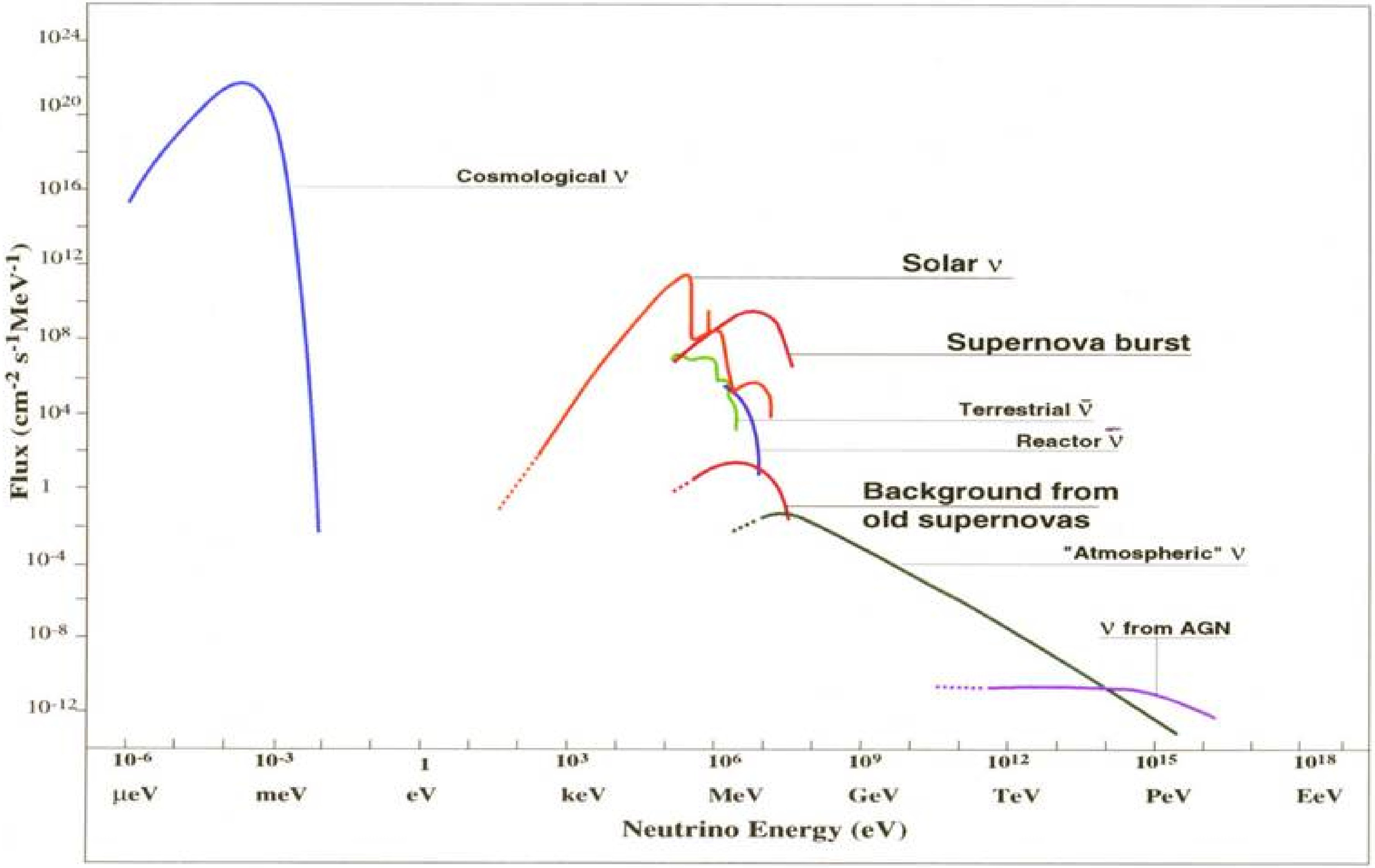
Lorentz Invariance Violation

Special Relativity Violation

etc…

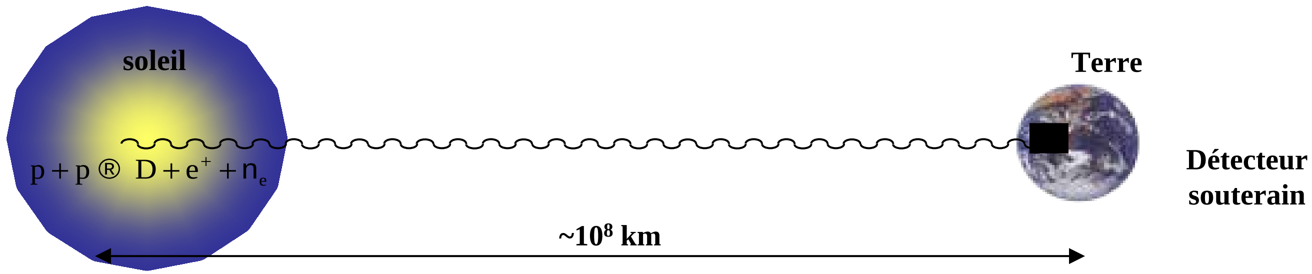
NEUTRINOS SOURCES

The overall neutrino spectrum

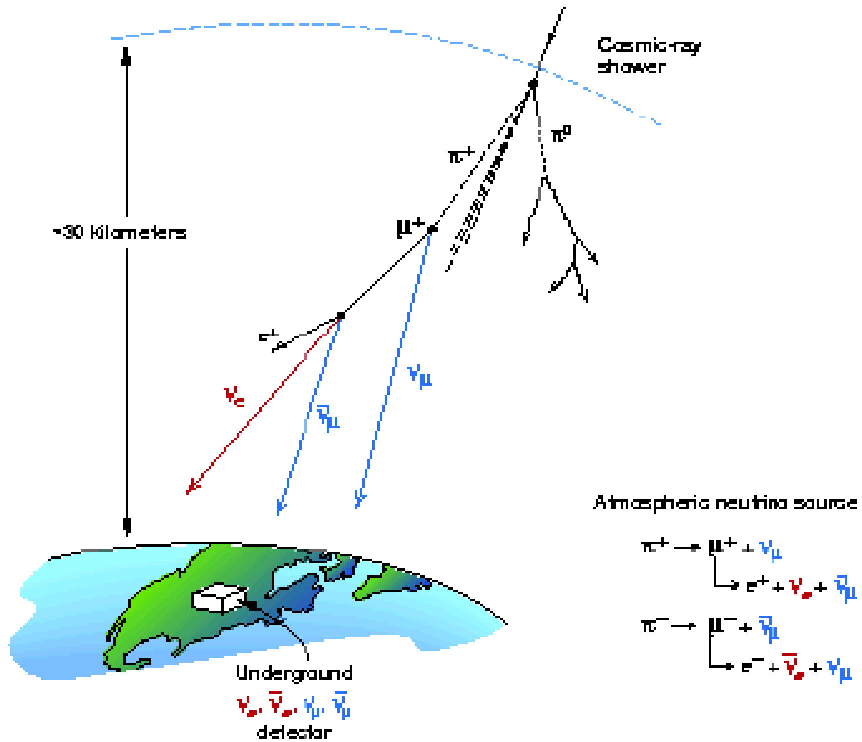


Natural Neutrinos Sources

Solar
Neutrinos



Atmospheric
Neutrinos



Also,
Super Novae (SN1987A),
Neutrinos de beam-dump cosmiques (AGN, GRB...),
Neutrinos cosmogéniques, Neutrinos reliques du BigBang...

Neutrino Production in the Sun

Light Element Fusion Reactions

$p + p \rightarrow {}^2H + e^+ + \nu_e$ (99.75%)

$p + e^- + p \rightarrow {}^2H + \nu_e$ (0.25%)

$^2H + p \rightarrow {}^3He + \gamma$

85%: $^3He + {}^3He \rightarrow {}^4He + 2p$

~15%: $^3He + {}^4He \rightarrow {}^7Be + \gamma$

~10^{-5}%: $^3He + p \rightarrow {}^4He + e^+ + \nu_e$

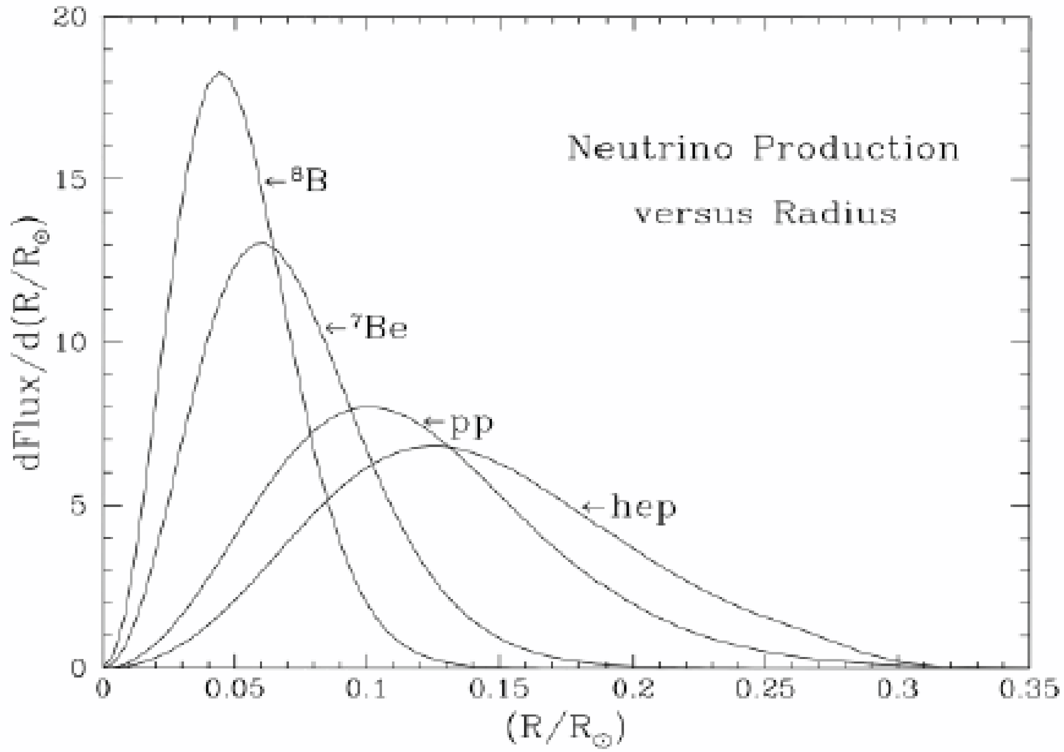
15.07%: $^7Be + e^- \rightarrow {}^7Li + \gamma + \nu_e$

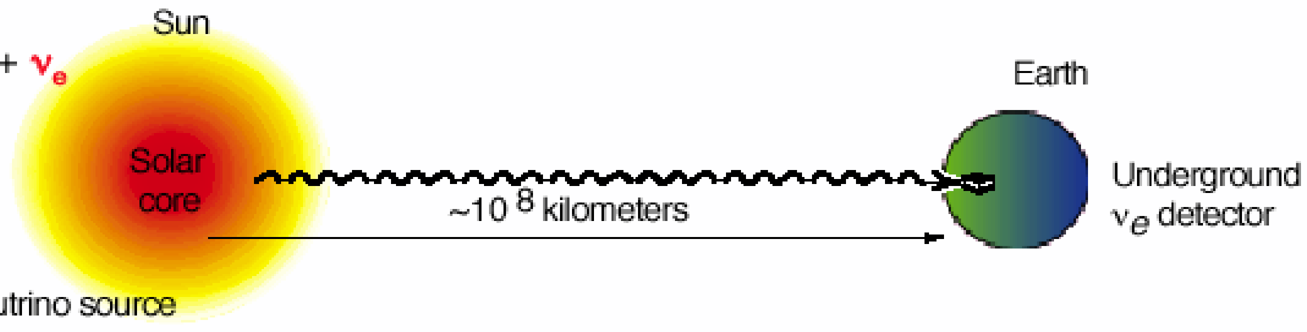
$^7Li + p \rightarrow \alpha + \alpha$

0.02%: $^7Be + p \rightarrow {}^8B + \gamma$

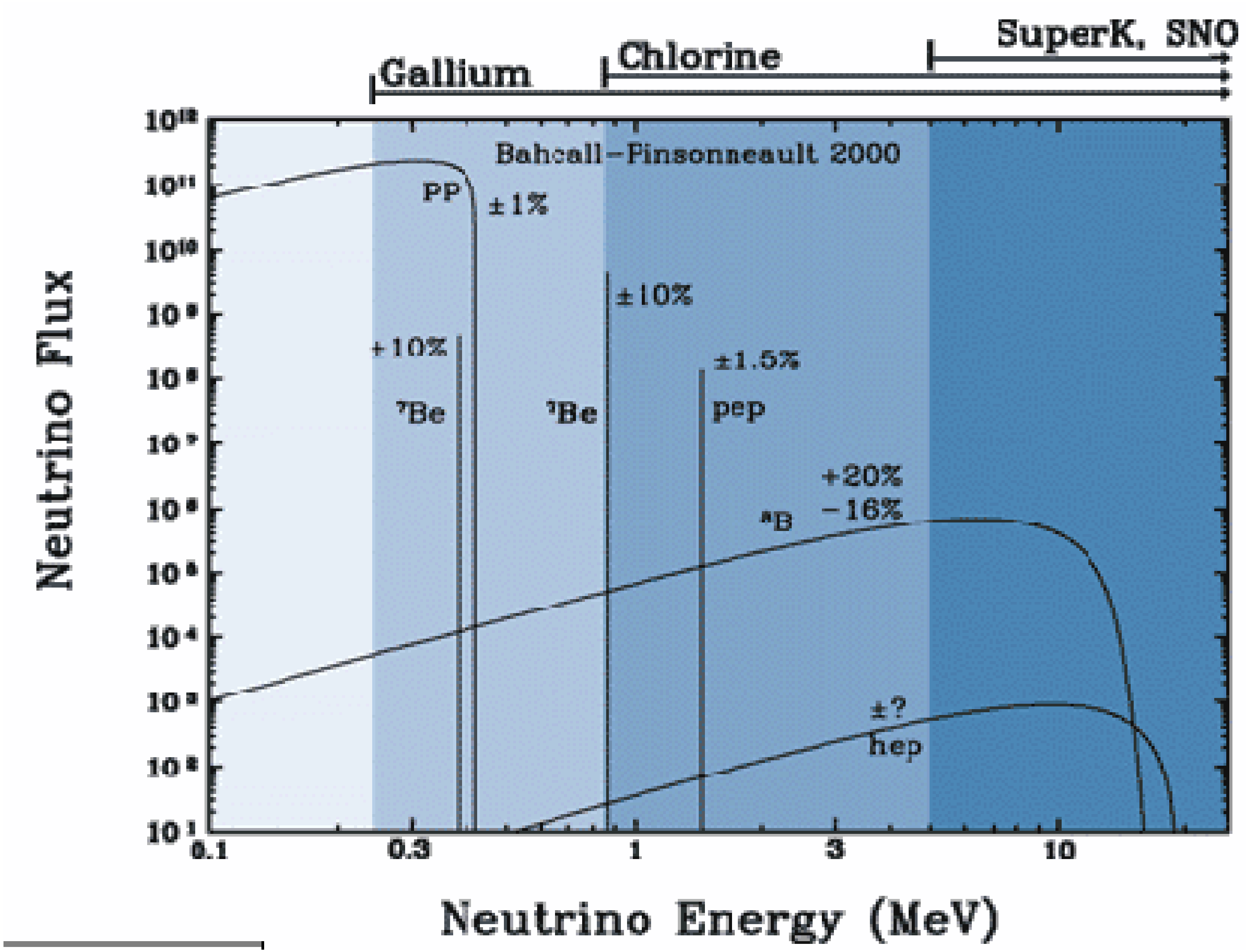
$^8B \rightarrow {}^8Be^* + e^+ + \nu_e$

Neutrino Production Radius





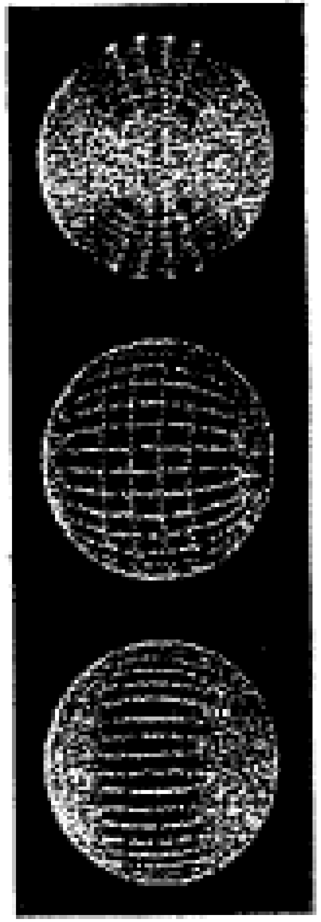
Karsten Heeger, LBNL


SuperK, SNO
Chlorine
Gallium
Bahcall-Pinsonneault 2000
Neutrino Flux
PP
±1%
±10%
+10%
±1.5%
⁷Be
⁷Be
pep
+20%
−16%
⁸B
±?
hep
0.1
0.3
1
3
10
Neutrino Energy (MeV)

Solar Standard Model

- *NEW* 39 experts agree on cross secs & systematics!

  Rev. Mod. Physics, Oct 1998, astro-ph/9805121



- How well do we know T?

  $\phi(pp) \sim T^{-1.2}$, $\phi(Be) \sim T^{8}$, $\phi(B) \sim T^{18}$

  $\rightarrow$ Helioseismology as a test...

  - The sun is a resonant cavity
    frequencies depend on $U = P/\rho$.
  - The SSM predicts U
    Measured/SSM $U(r)$ agree $< 1\%$!
  - Temperature Uncertainty:
    SSM alone: $\Delta T/T = \pm 2.7\%$
    SSM + Helio: $\Delta T/T = \pm 1.4\%$

SNe neutrinos

Normal stellar situation: the fusion thermal and radiation pressure compensate the gravitation pressure.

During the collapse, the gravitation binding energy (¼ 3´ 10^{53} erg) cannot escape in an other form than neutrino-antineutrino pairs.

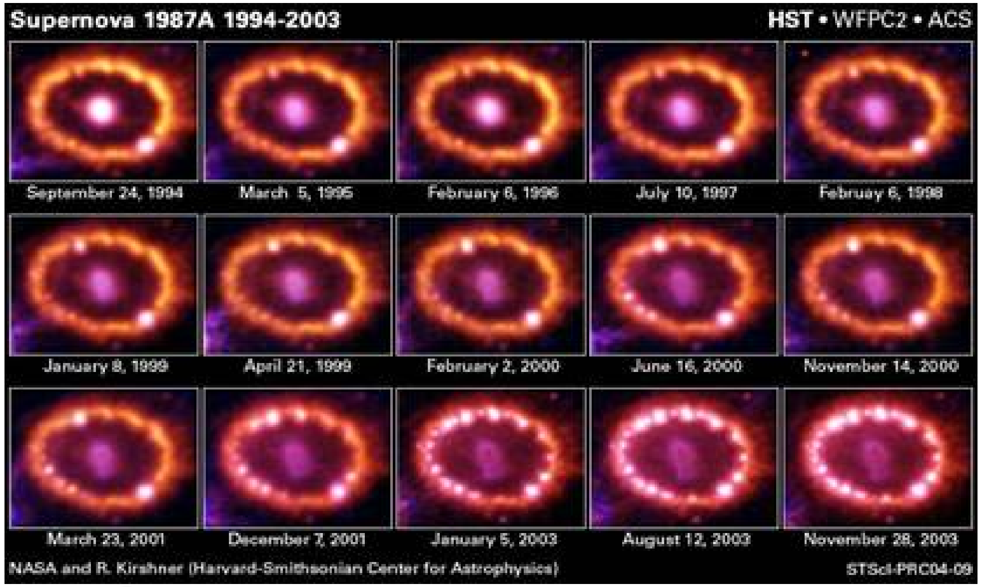
99% of the energy in the form of neutrinos
1% in the form of kinetic energy
0.01% of the energy in optical photons.

The neutron star is opaque to neutrinos. The diffusion and escape time is of the order of 1 second.

$$< E n_e >» 11 \text{ MeV}$$

$$< E \bar{n}_e >» 16 \text{ MeV}$$

SN1987a



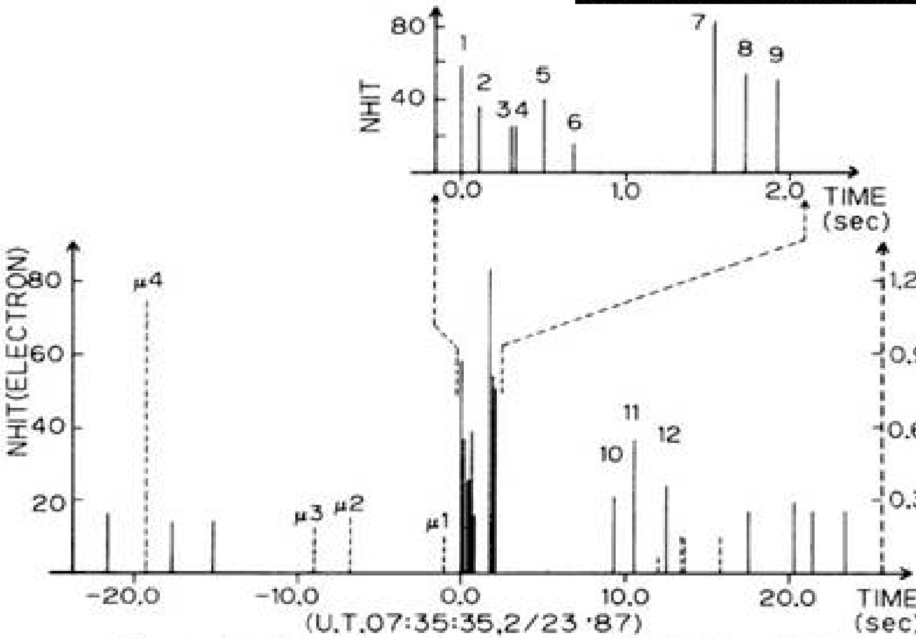
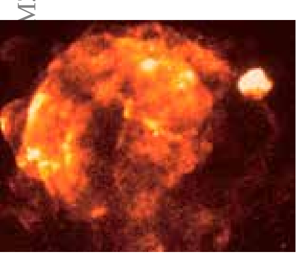


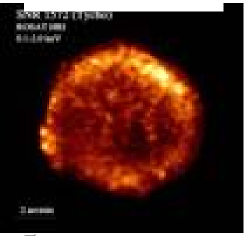
FIG. 2. The time sequence of events in a 45-sec interval centered on 07:35:35 UT, 23 February 1987. The vertical height of each line represents the relative energy of the event. Solid lines represent low-energy electron events in units of the number of hit PMT's, $N_{\rm hit}$ (left-hand scale). Dashed lines represent muon events in units of the number of photoelectrons (right-hand scale). Events $\mu 1-\mu 4$ are muon events which precede the electron burst at time zero. The upper right figure is the 0–2-sec time interval on an expanded scale.

Super Novae Remnants

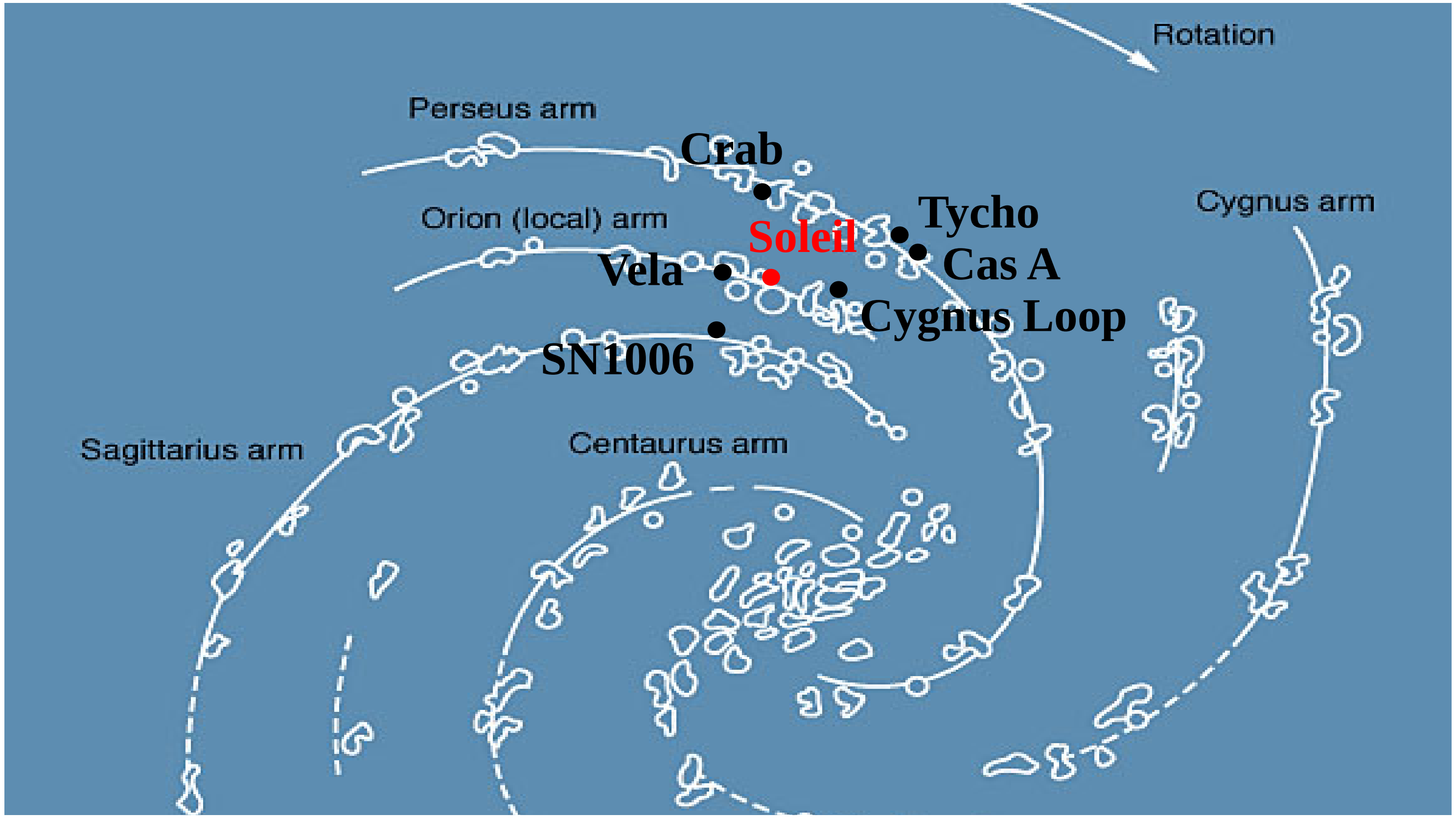
Vela

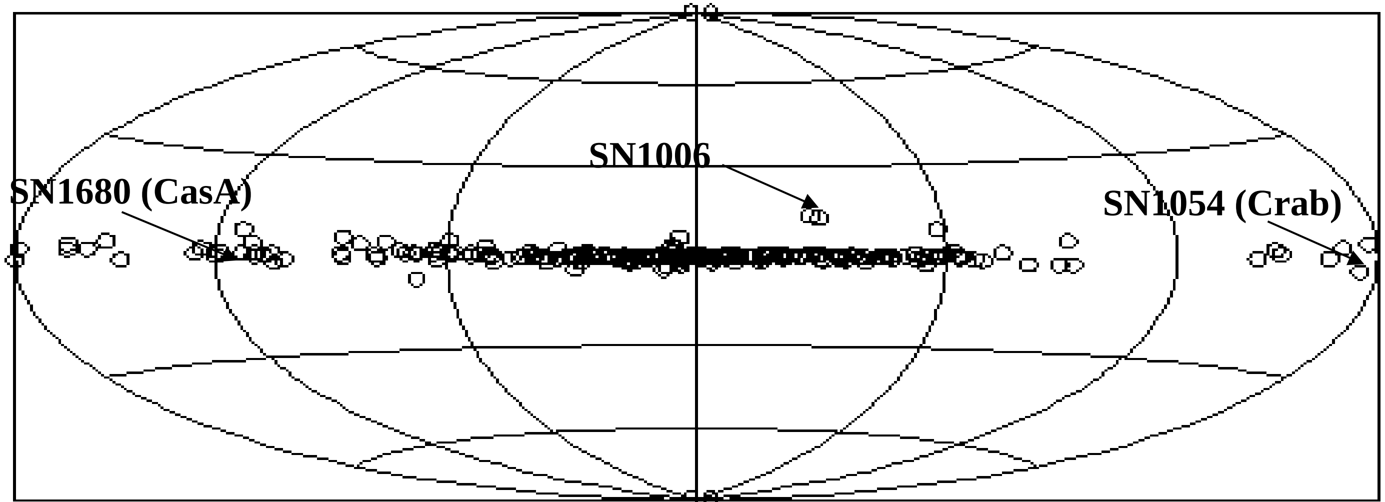


Tycho



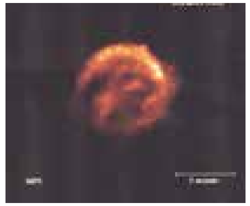
Cygnus



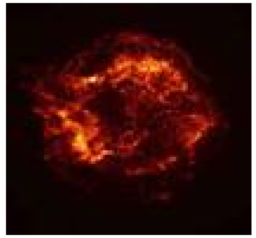


Crab

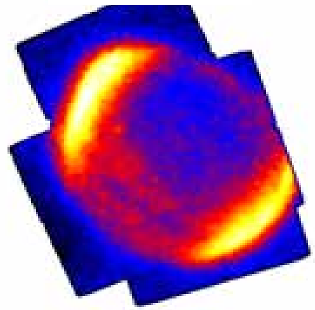
Kepler



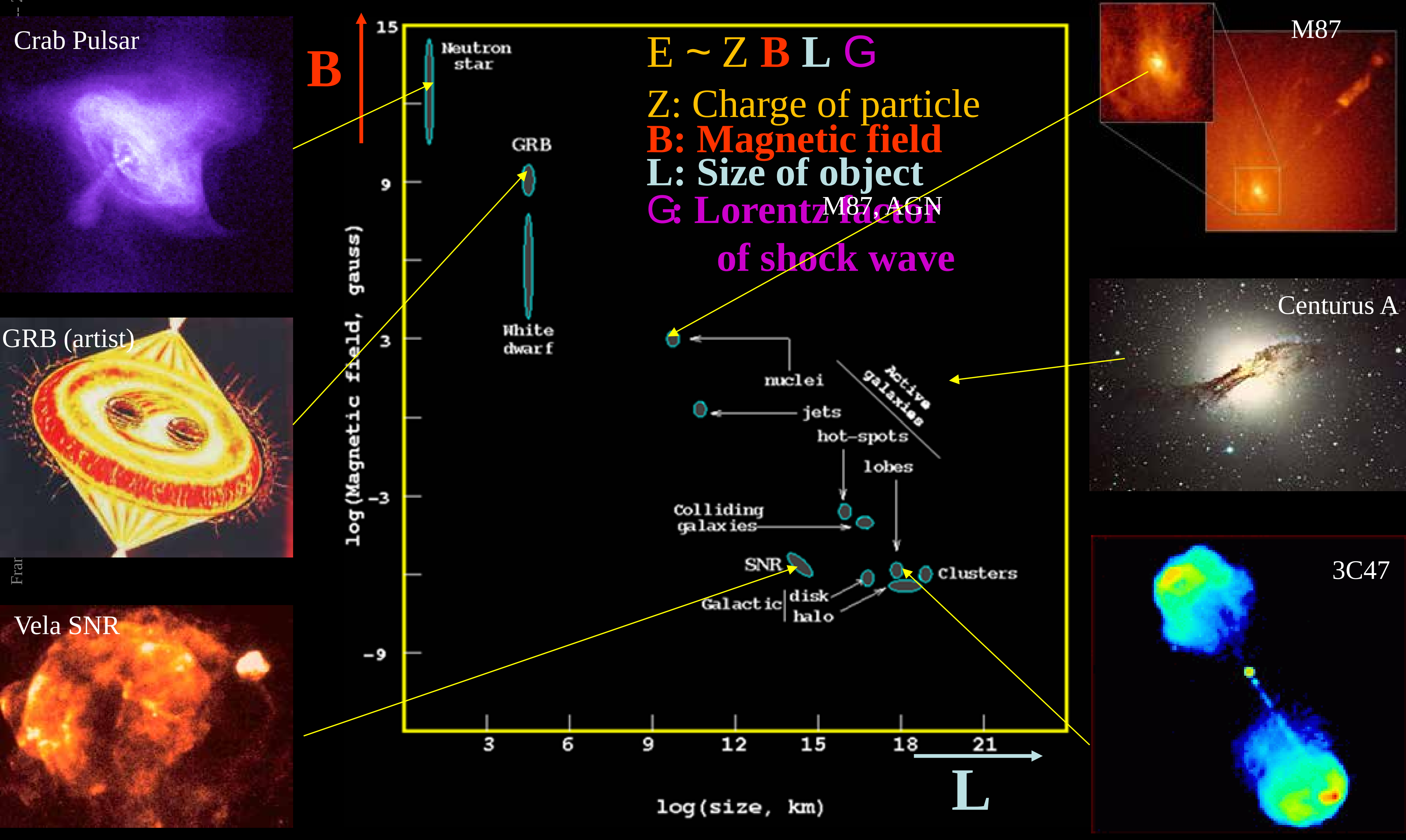
Cas A



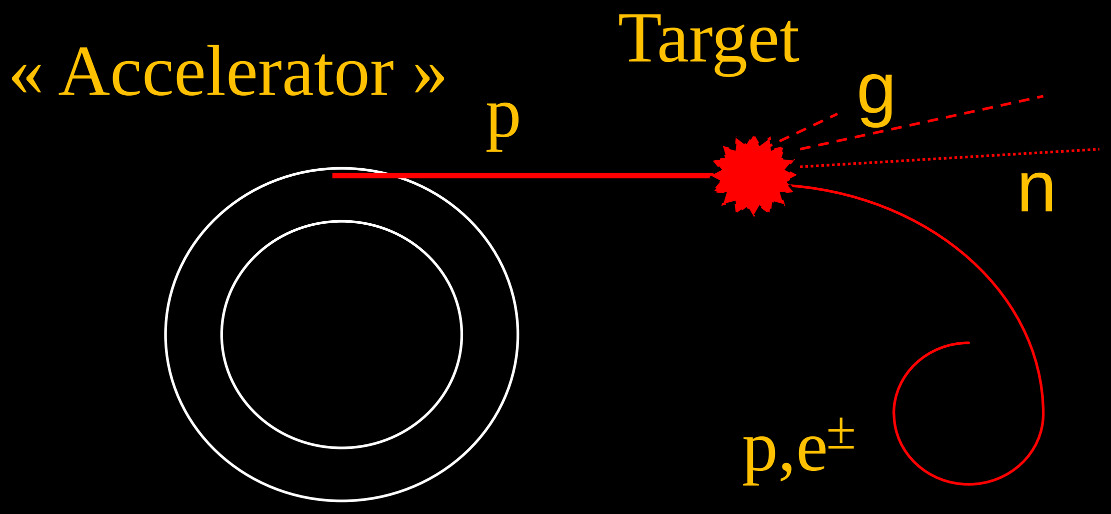
SN1006

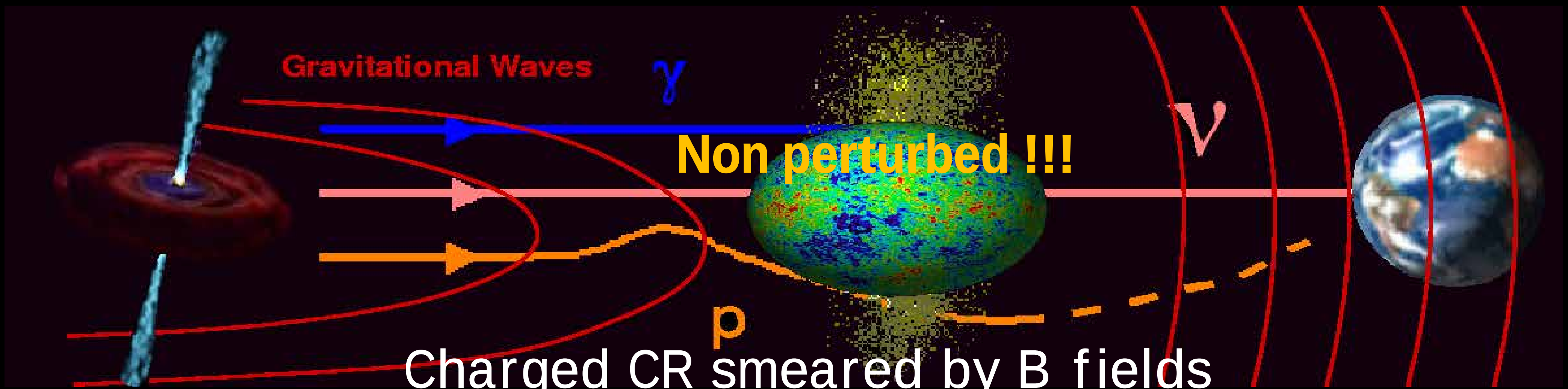


Cosmic Accelerators: (Hillas Plot)

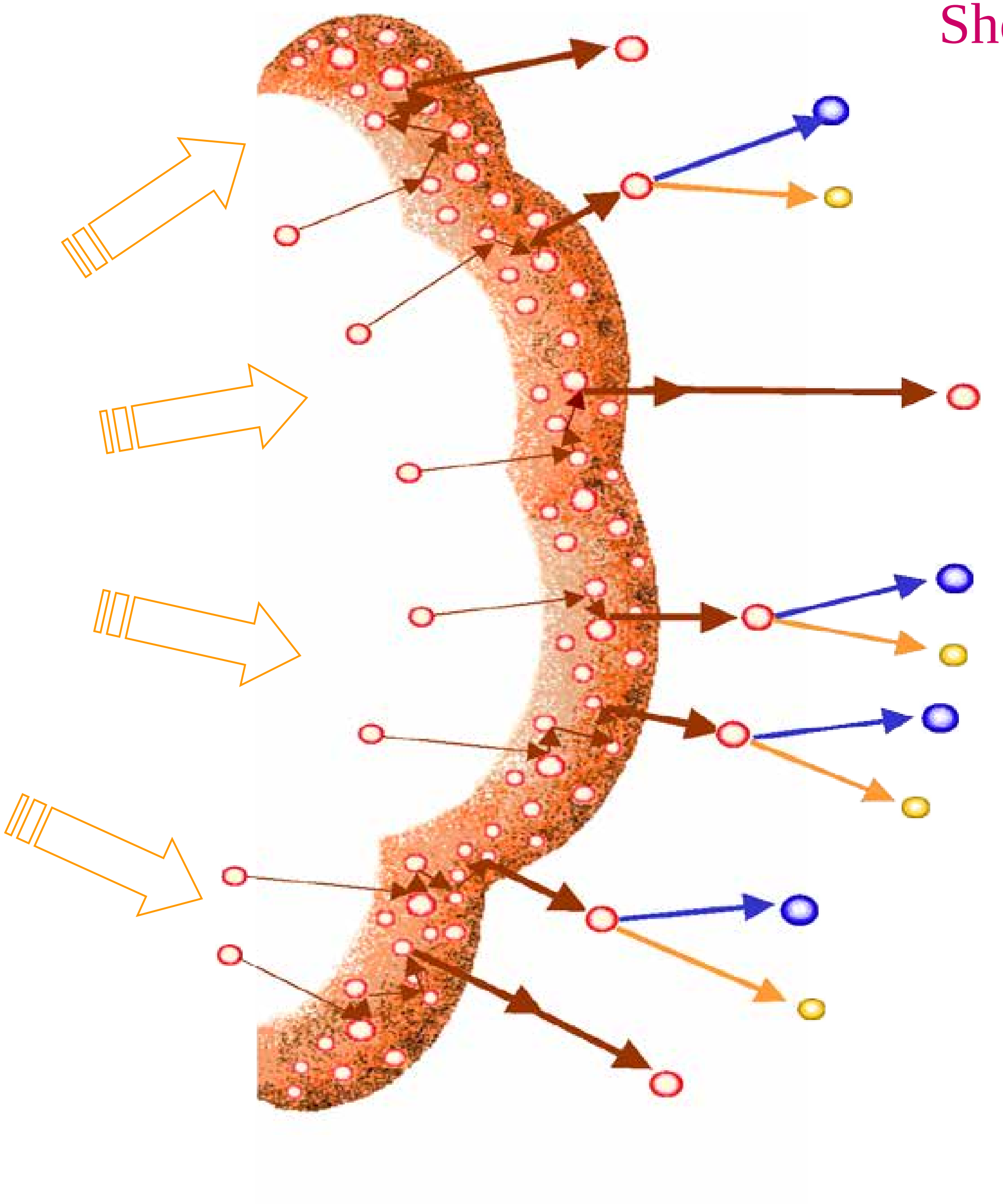


Cosmic Neutrino Beam





Fermi acceleration and UHE neutrino production



Shock wave proton of nuclei acceleration:

p/N interact with the ISM and produce pionic cascades

$p/A + p/g$ ® p^0 + $p^{\pm}$ + ...

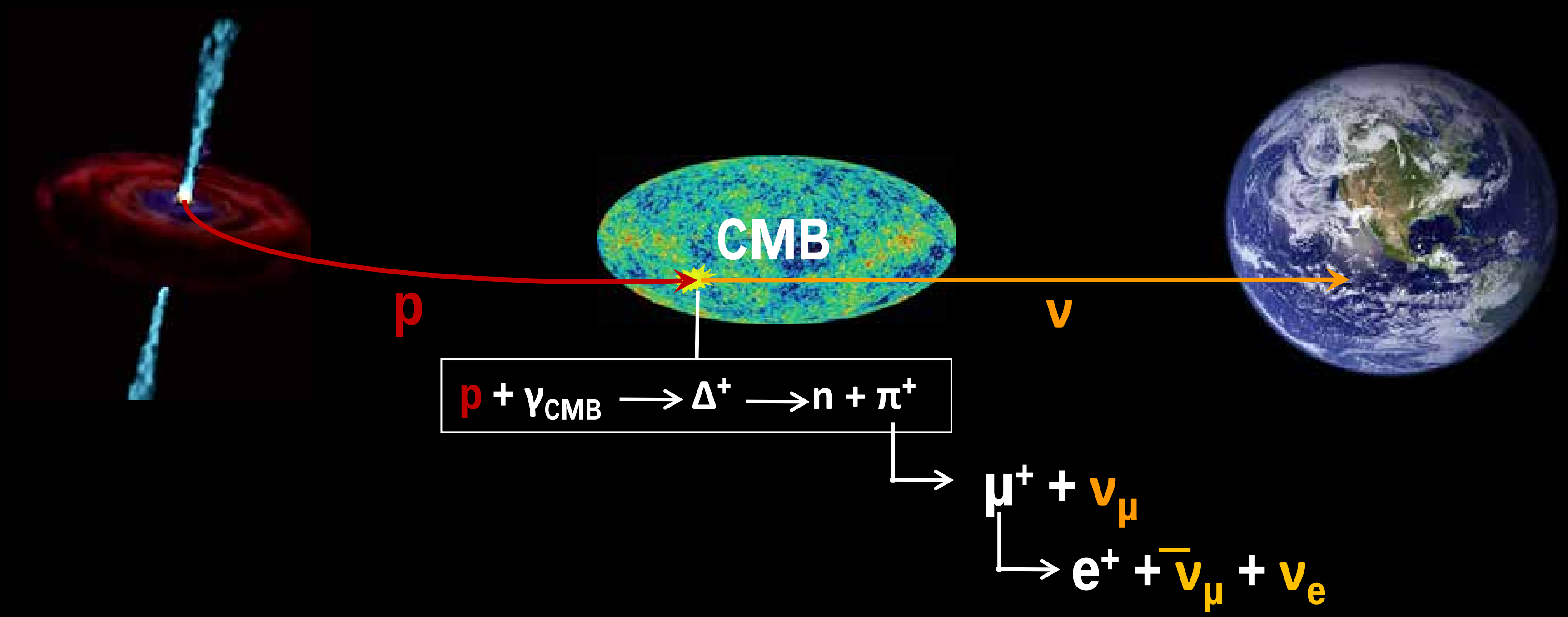
g g n_m m

n_m n_e e

COSMOLOGICAL SOURCES

"Cosmogenic" or "GZK" Neutrinos

- Produced by the propagation of UHECR in the IGM.
- Main assumptions:
  - UHECR are protons or nuclei (makes a big difference!)
  - The sources distribution is following a given redshift distribution
  - Spectra are known and flux are generic
  - The target density is well known (CMB photons)
- This flux can be predicted in a relatively robust manner.



Cosmogenic or "GZK" Neutrinos

Flux attendus

I - Sources

- Composition at the source
- Spectral index + E_{max}
- Source density + cosmological evol.

Many uncertainties

II - Fonds diffus

- Backgrounds (CMB, CIB...)
- Photon spectrum evol vs z

CMB very well known.
CIB and other bg less known

III - Interaction pγ

- Cross Section
- Secondaries energy distribution
- Energy Distribution

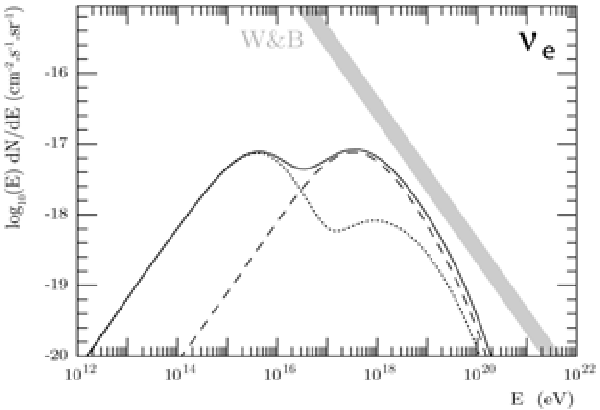
Very well known

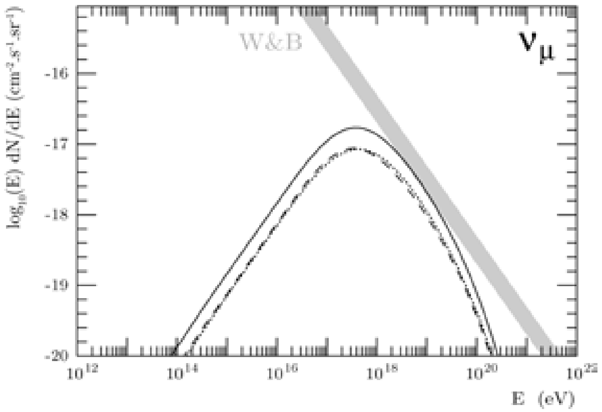
I + II + III è

$\nu_\mu:\nu_e:\nu_\tau$ = 2:1:0 (source)

oscillations

1:1:1 (earth)

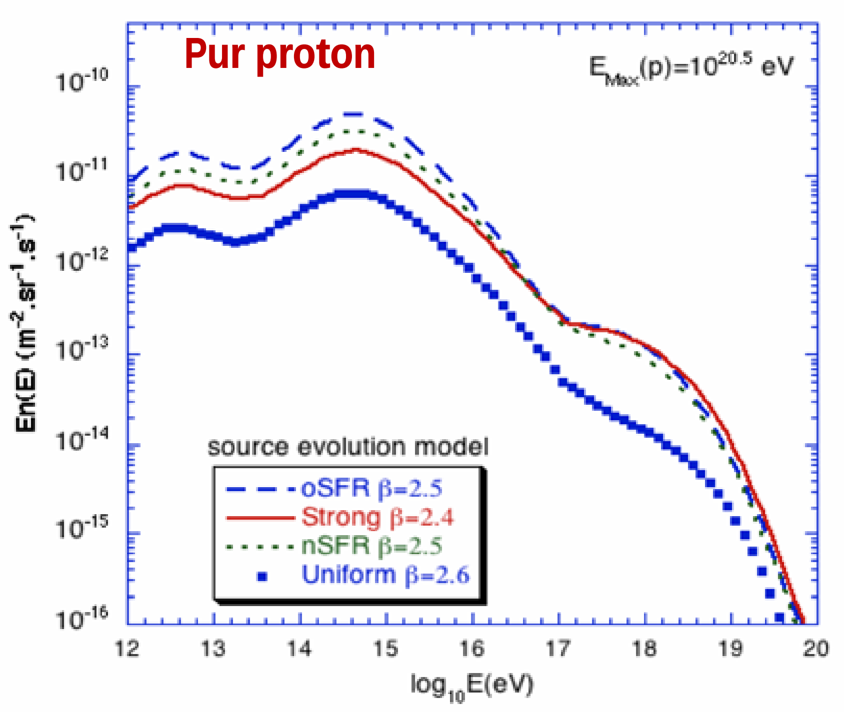


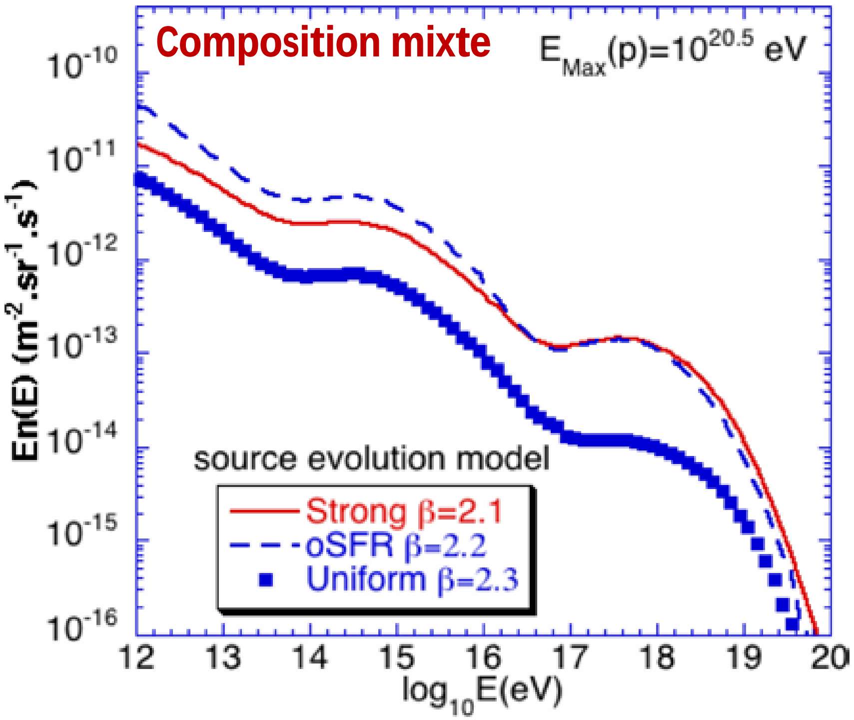


Flux of "GZK" neutrinos

Influence of

- Composition of UHECR,
- Sources distribution and evolution.



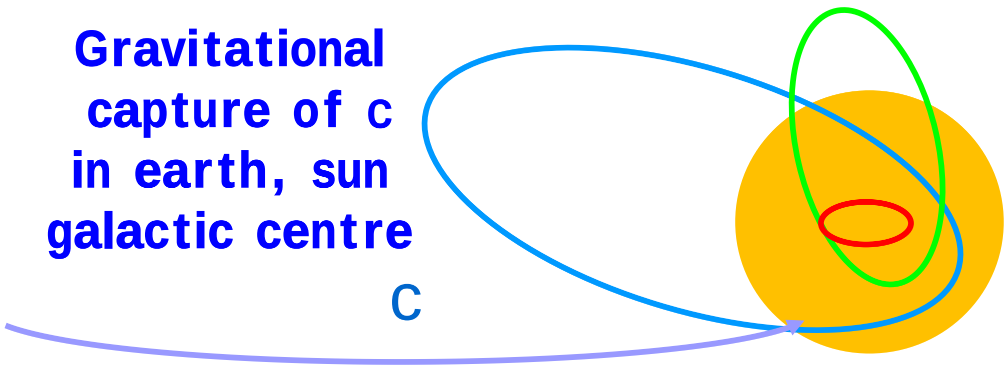


Dark Matter as a source of neutrinos

Annihilation in Halo, Earth, Sun or Galactic Centre

Signature	Experiment
Halo **Positron, Antiproton** **Gamma rays** $\chi\chi \rightarrow Z\gamma, \gamma\gamma$	**BESS, CAPRICE, AMS, ..** **HESS, GLAST, MILAGRO,….**
Earth, Sun, GC **Neutrino** $\chi\chi \rightarrow WW, ff$ $W, f \rightarrow \nu X$	**SuperK, Baksan, IMB, MACRO** **AMANDA, ANTARES, Baikal, …**

Gravitational capture of χ in earth, sun galactic centre



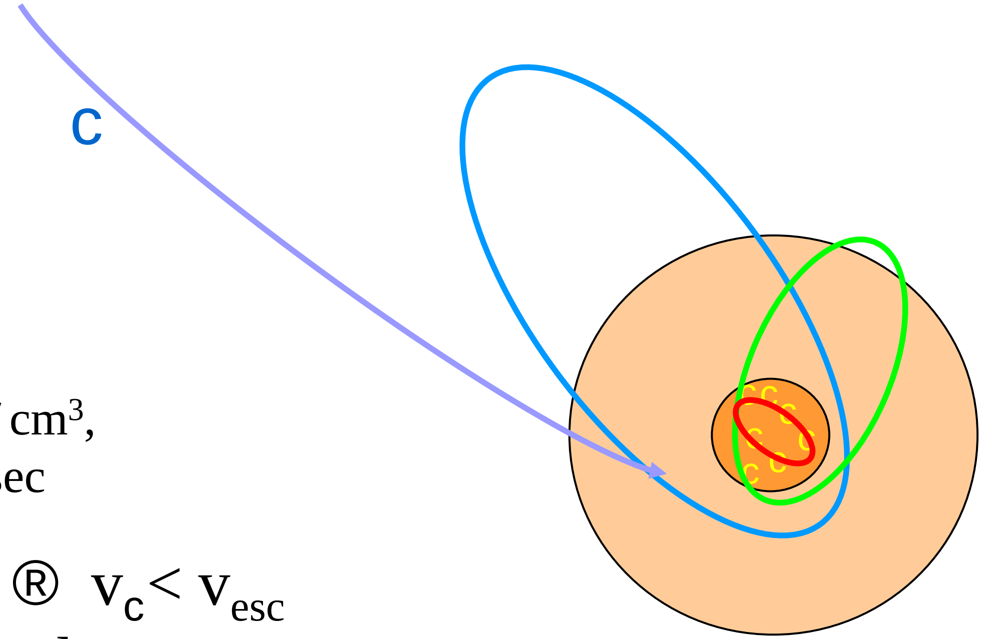
WIMP looses energy by elastic interaction

⇒ if $v < v_{escape}$, capture

capture + annihilation balance

⇒ constant density in core

Dark Matter as a source of neutrinos



DM halo

r $_c$ ~ 0.3 Gev/ cm^3,
v$_c$~300 km/sec

Collisions ® v$_c$< v$_{esc}$
Gravitational capture
in the center
gravitational wells

$$\chi + \chi \rightarrow WW \quad bb \quad \ldots$$
$$\qquad\qquad\quad \downarrow \qquad \downarrow$$
$$\qquad\qquad\quad \mu\nu \qquad c\mu\nu$$
$$E_\nu \sim 0.3 - 0.5 \times m_\chi$$

Þ Indirect search for Neutralinos by neutrino telescopes

Propagation

The general problematic

- Thermal speeds ! RCUHE (few 10^{20} eV)

Produce them

- From top to bottom (decay...)
- From bottom to top (acceleration)

Preserve them

- Energy losses (Synch., IC, p, pairs...)
- Destruction (photo-dissociation...)
- Escape probabilities

Propagate them

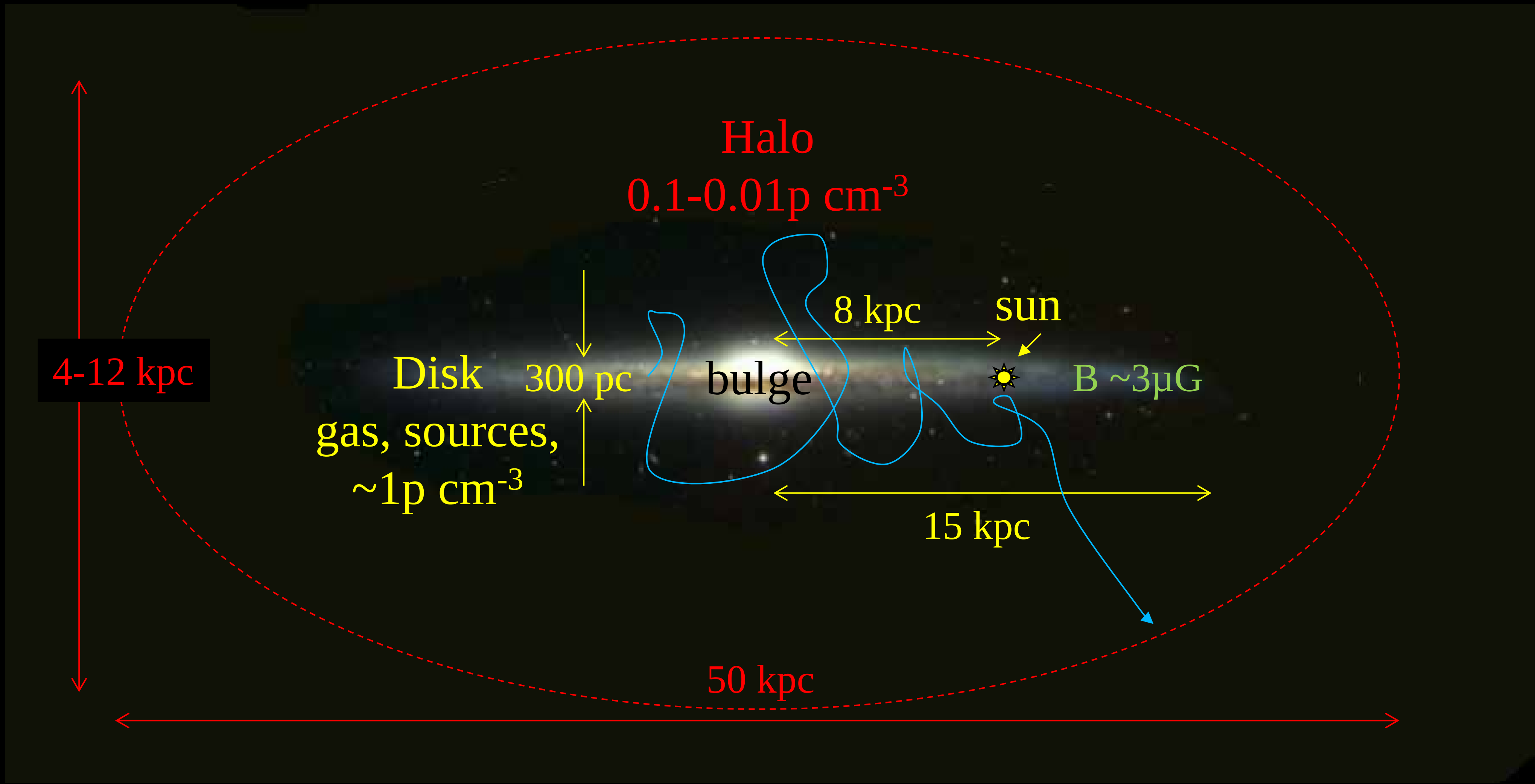
- Propagation in ISM and IGM (mag fields: deflection, confinement...)
- Re-acceleration

Detect them

- Balloons, satellites...
- Air showers...
  - Cherenkov telescopes
  - Surface & Fluorescence Detectors

Dimensions of the Milky Way

1 pc ¼ 3 l.y. ¼ 3 £ 10^{16} m

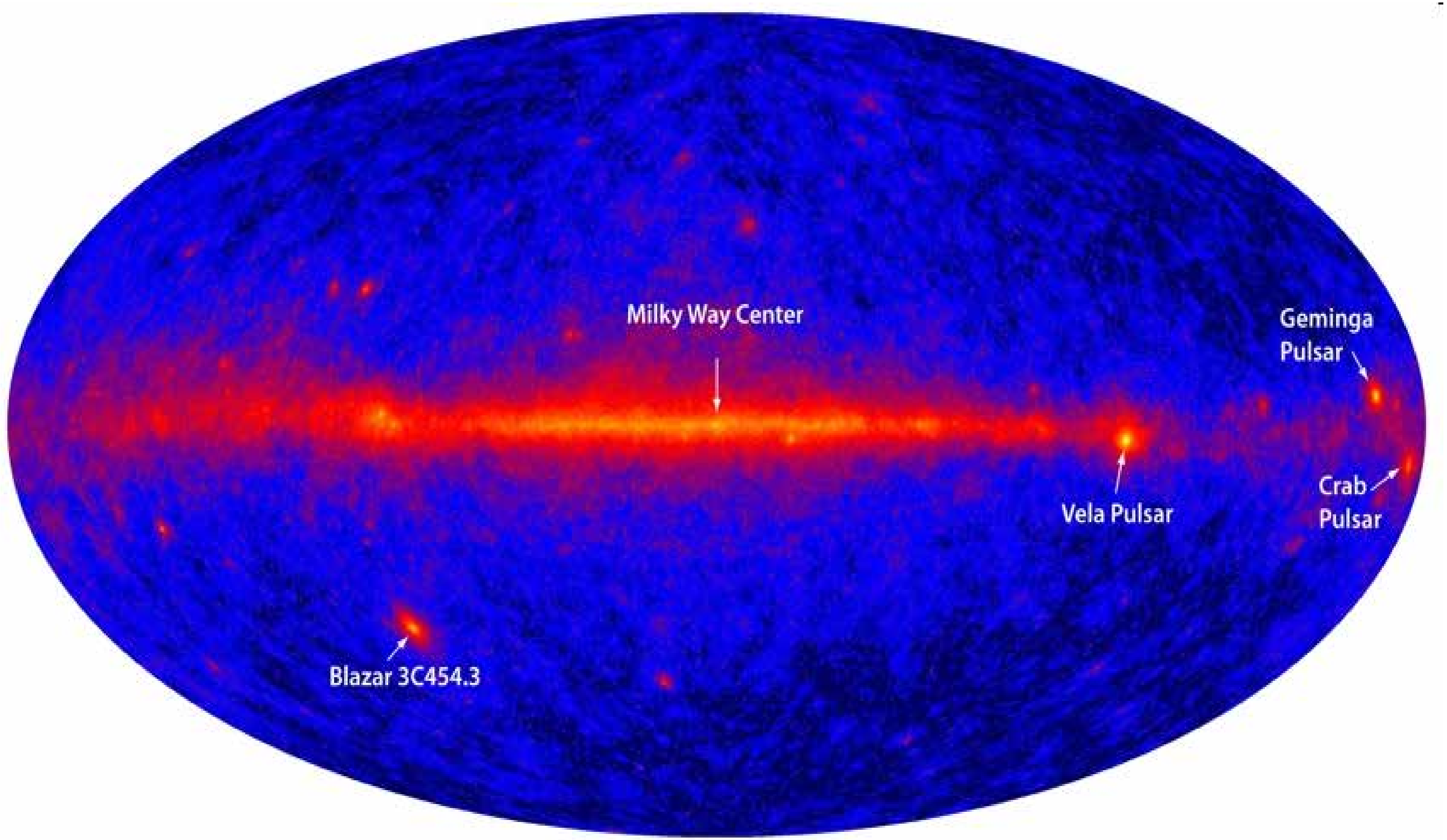


F.Montanet Experimental Astroparticle Physics ESIPAP 2014

Propagation des RC dans la Galaxie : Leaky box model

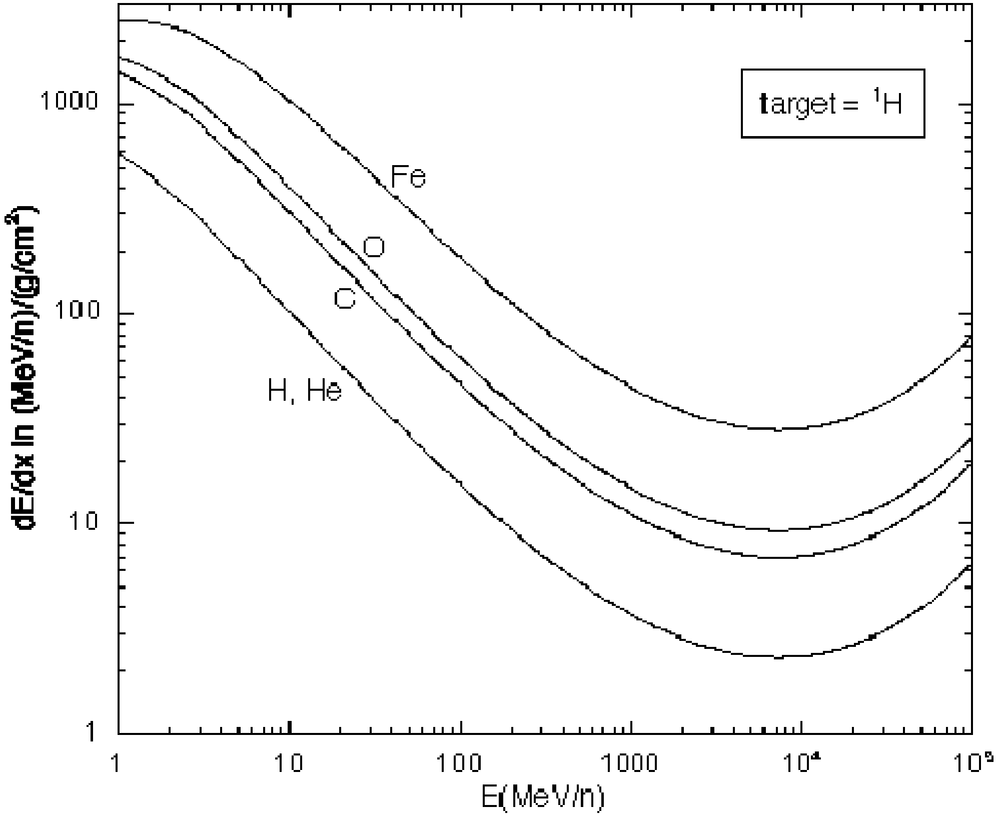
A thick target

- Diffuse gamma-ray emission from galactic CR interaction with matter (mostly molecular H clouds).

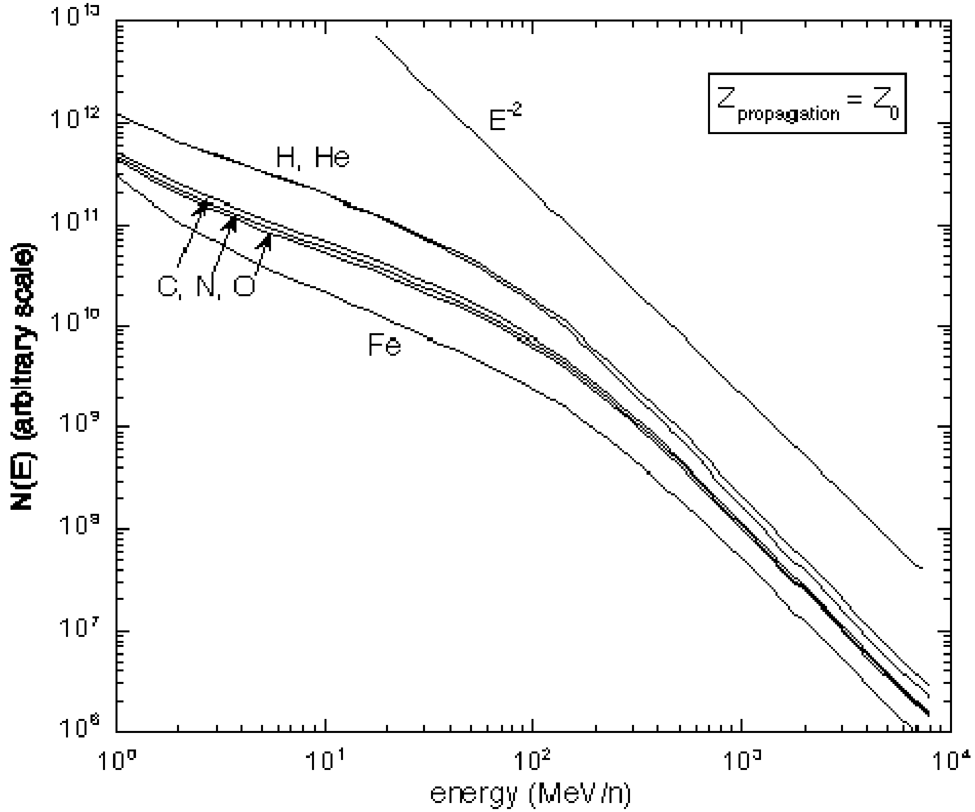


Cosmic rays transport

- Propagation in the interstellar medium



Energy loss: ionization, Coulombian interactions



Propagated spectra ionization losses only (thick target)

Grammage

- Column density or quantity of matter traversed by the CR from is production site to earth (in kg ¢ m^{-2} or g ¢ cm^{-2})
- Given the diffusion time (known from cosmic clocks see below) the measurement of grammage allows the understanding of the diffusion extension zone.
- The ratio secondary/primary allows estimating the grammage traversed:

$$\frac{dN_S}{dx} = -\frac{\sigma_P}{m} N_P$$

$$\text{donc } N_S = N_P \exp -\frac{\sigma_P}{m} x$$

$$\text{et } x = -\frac{m}{\sigma_P} \log(S/P)$$

$$\text{B/C} \approx 35\% \Rightarrow x = -\frac{m}{\sigma_P} \log(B/C) \approx 60 kg.m^{-2}$$

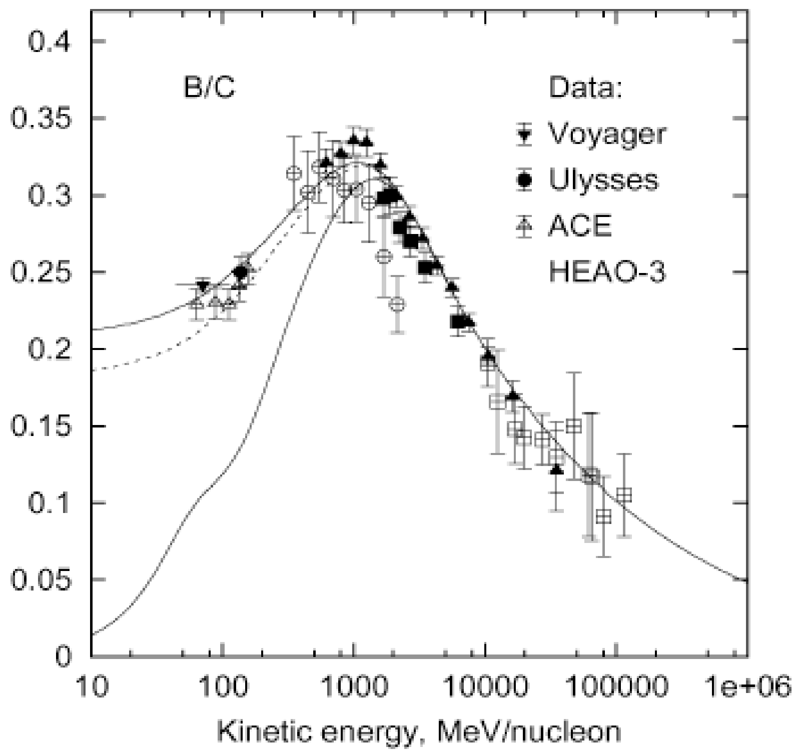
$$\text{if } Br(C + P \to B + X) \approx 100\%$$

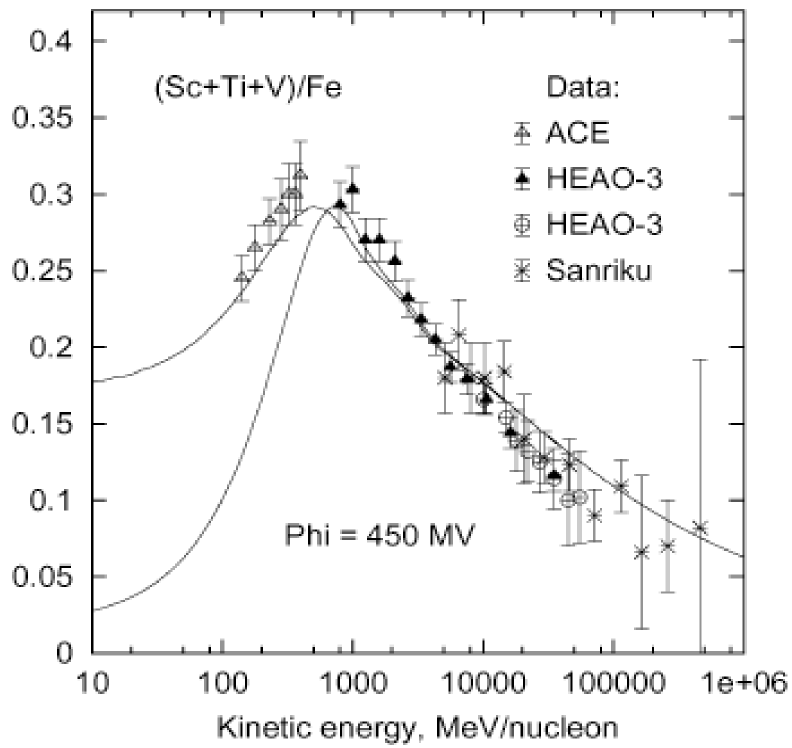
Secondary / Primary ratio

The grammage depend on the parent nucleus:

- (Li+Be+B) / (C+N+O) Þ mean grammage of 50 $kg.m^{-2}$
- (Sc+Ti+V)/Fe Þ mean grammage of 20 $kg.m^{-2}$

and Primary/Secondary ratio (thus grammage) depends on the energy as well:





Secondary / Primary ratio

- A complete CR transport model
  - **The secondary to primary ratio can be expressed by:**

$$\frac{N_S}{\tau_{esc}} + \frac{N_S}{\tau_{spallation}} = \frac{N_P}{\tau_{P\rightarrow S}}$$

$$\Rightarrow \frac{N_S}{\tau_{esc}} + n\beta c\sigma_S N_S = n\beta c\sigma_{P\rightarrow S} N_P$$

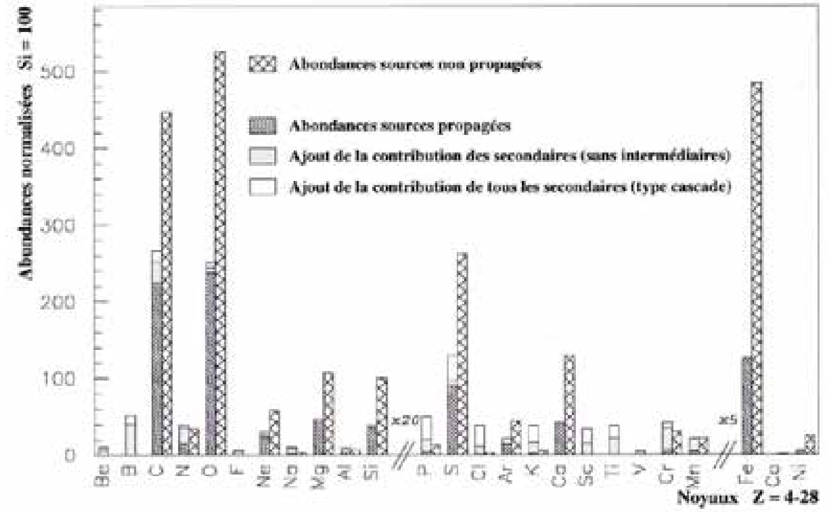
$$\frac{N_S}{N_P} = \frac{\sigma_{P\rightarrow S}}{\sigma_S + 1/\lambda_{esc}}$$

with $\lambda_{esc} = n\beta c\tau_{esc}$

Cosmic clocks

Unstable nuclei with lifetimes comparable to the escape time $T_{1/2} \approx \tau^{esc}$ can be used as cosmic clocks.
The ratio unstable/stable isotope helps desantangling density and escape time.

$$^{10}Be \rightarrow \tau = 2.17 Myr$$
$$^{26}Al \rightarrow \tau = 1.31 Myr$$
$$^{36}Cr \rightarrow \tau = 0.44 Myr$$



$$\frac{N_j}{\tau_{esc}} + \frac{N_j}{\tau_{rad}} + \frac{N_j}{\tau_{spallation}} = Q_j + \sum_{k>j} \frac{N_j}{\tau_{k\rightarrow j}}$$

Si $\tau_{rad} \ll \tau_e$ et $\tau_{rad} \ll \tau_{spallation}$:

Measure isotopic ratio →

$$\frac{N_{rad}}{N_{stable}} = \frac{\tau_{rad}}{\tau_{esc}} + \frac{\tau_{rad}}{\tau_{spallation}}$$

→ Estimate escape time.

On gets $\tau_e \approx 20 Myr$

Cosmic clocks and halo size

- Radioactive decay:

$$\frac{N_j}{\tau_{esc}} + \frac{N_j}{\tau_{rad}} + \frac{N_j}{\tau_{spallation}} = Q_j + \sum_{k>j} \frac{N_j}{\tau_{k \to j}}$$

Measure isotopic ratios →

$$\frac{N_{rad}}{N_{stable}} = \frac{\tau_{rad}}{\tau_{esc}} + \frac{\tau_{rad}}{\tau_{spallation}}$$

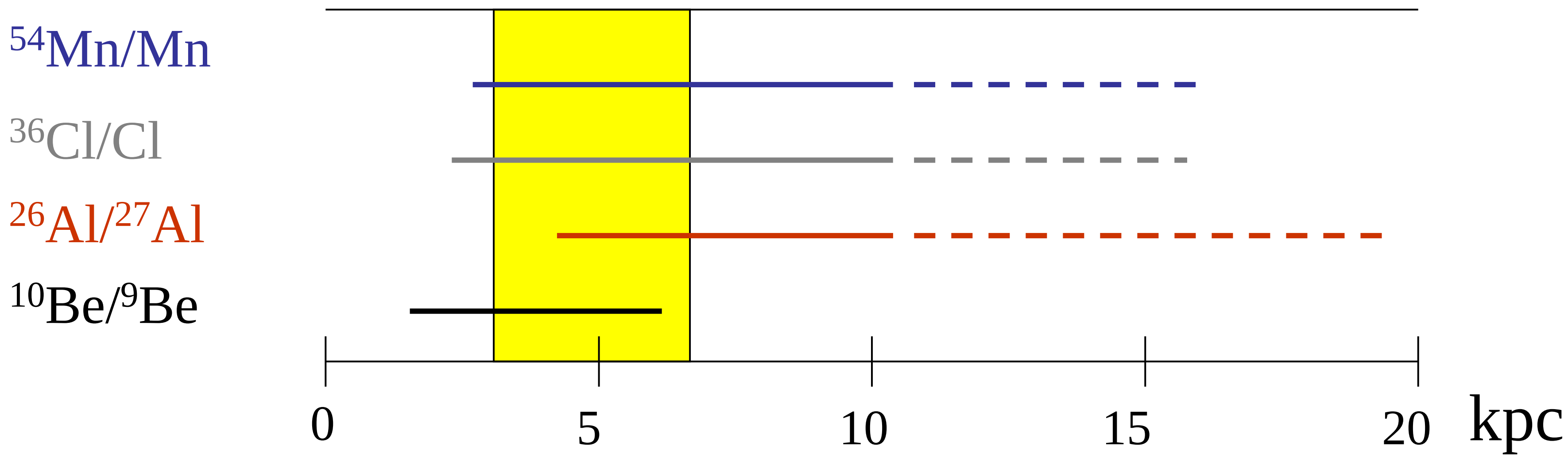
→ Estimate escape time

- ^{12}C + H ® ^{9}Be (stable secondary nucleus)
  ^{12}C + H ® ^{10}Be (unstable secondary nucleus: ~ 4´ 10^{8} years)
- The ratio ^{10}Be / ^{9}Be depends on secondaries **history** (and on cross sections).
  - Link between quantity of matter traversed and diffusion time.

Cosmic clocks and halo size

- ^{12}C + H ® ^{9}Be (stable secondary nucleus)
  ^{12}C + H ® ^{10}Be (unstable secondary nucleus: ~ 4´ 10^8 years)
- The ratio ^{10}Be / ^{9}Be depends on secondaries **history** (and on cross sections).
  - Link between quantity of matter traversed and diffusion time.
- Diffusion parameters adjustments (excursion in the less dense galactic halo)

  Þ determination of the CR confinement zone



Confinement and escape

- The average measured grammage is $x = 50 kg.m^{-2}$
- Associated lengths:

$$\begin{aligned} \lambda_{esc} &= x/\rho \approx 750 kpc, \\ \text{with} \quad \rho &= 1.4\, n_H\, m_p \approx 2.2 \times 10^{-21} kg.m^{-3} \end{aligned}$$

- $\lambda_{esc} \gg R = 20 kpc$ $\Rightarrow$ CR are confined
- $\lambda_{esc} \ll \lambda_{pp} = (n_H\, \sigma)^{-1} \approx 6 Mpc$ $\Rightarrow$ CR can escape
- Long lived radioactive secondaries (cosmic clocks) indicate $\tau_{esc} \approx 20 Myr$
- Average density scanned by CR:

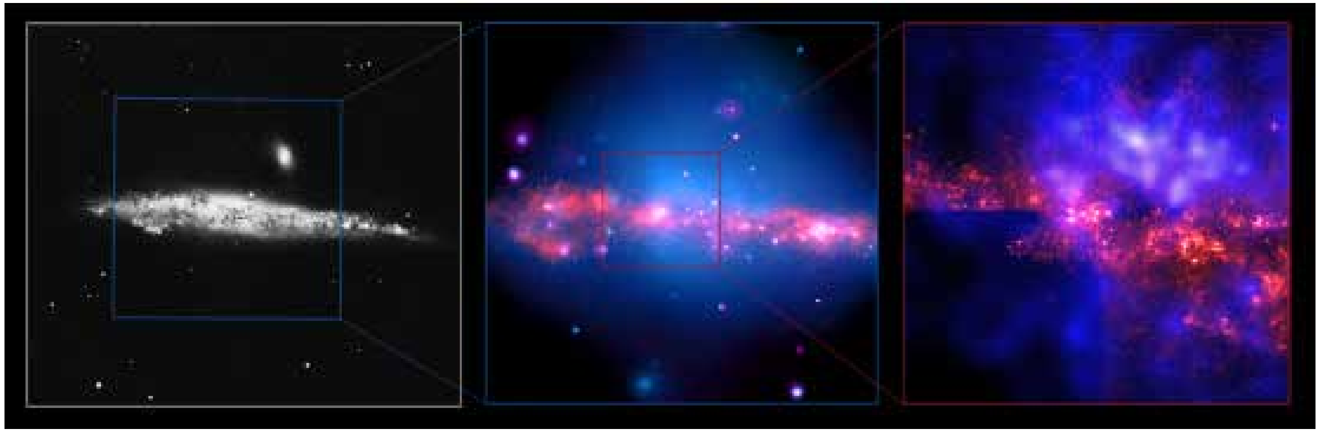
$$n_H = \lambda_{esc}/c\, \tau_{esc}\, m_p \approx 0.3 cm^{-3} < n_{disk} = 1 cm^{-3}$$

$\Rightarrow$ CR diffuse in a thinner region: the Halo

Disk & Halo

- CR can wander out of the disk in a magnetized halo of hot ionized matter

$$T = 10^6\ K \quad \text{et}\ n = 10^{-3} cm^{-3}$$



- NGC 4631 galaxy and its halo of hot ionized matter emitting X-rays as seen by Chandra

The « leaky box » model

- Diffusive approximation...

$$\frac{\partial N_i}{\partial t} + \frac{\partial}{\partial E}[b(\vec{r},E)N_i(\vec{r},E,t)] = Q_i(\vec{r},E) - \frac{N_i}{\tau_{tot}(\vec{r},E)} + D(\vec{r},E)\nabla^2 N_i$$

flux in enegy space (losses + acceleration)

injection, "in flight" production

destruction, decay, escape

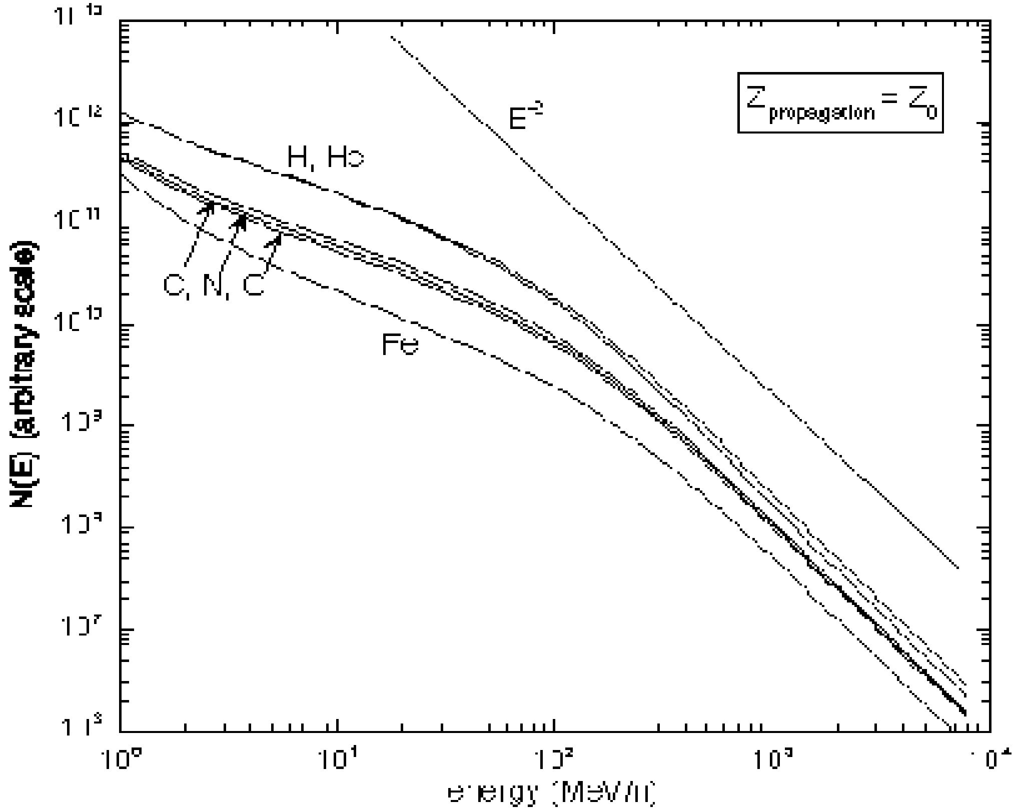
diffusion

- Re-acceleration...

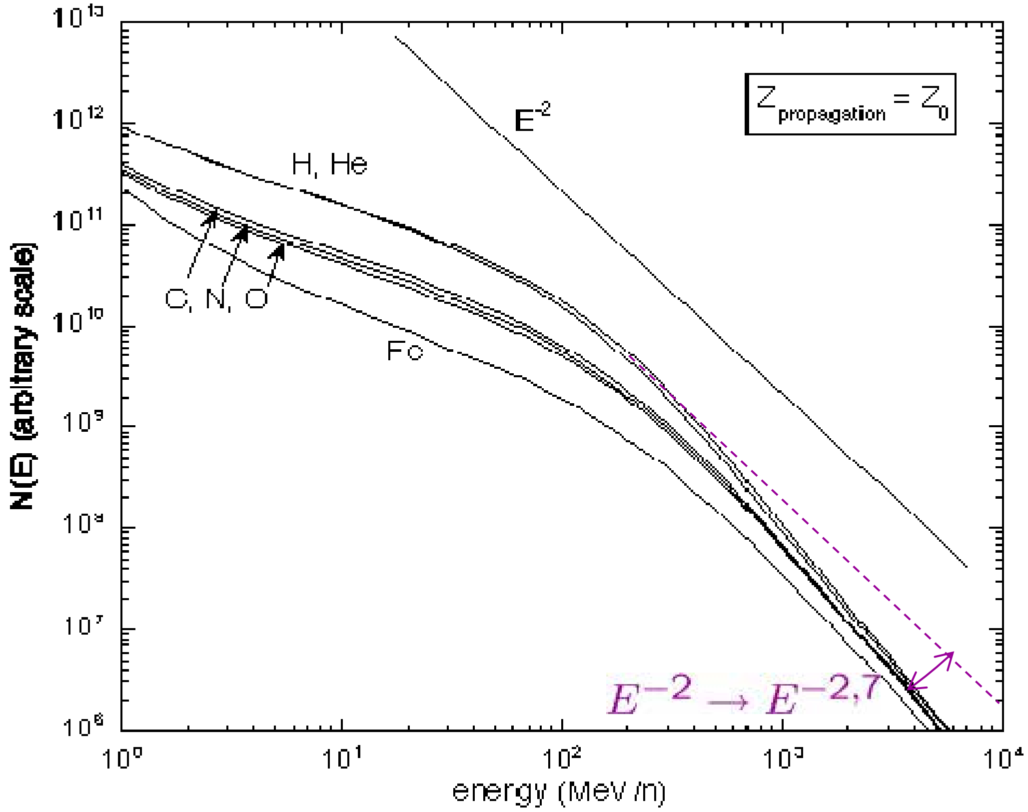
$$+\frac{1}{2}\frac{\partial^2}{\partial E^2}[c(\vec{r},E)N_i(\vec{r},E,t)]$$

Slope of the propagated spectrum

- Escape out of the confinement zone
  - Confinement (escape probability) decrease with E



Without escape
(thick target)



t_{conf} µ $E^{-0.7}$
® $E^{-2.7}$ spectrum

CR confinement

- Escape depends on E
  - Diffusion on magnetic inhomogeneities
  - When E k, r_g k thus interaction with inhomogeneities with larger wavelengths.
- ***D*(*E*)** is an increasing function

$$D = \beta D_0 \left(\frac{\rho}{\rho_0}\right)^x \text{ where } \rho \text{ is the particle rigidity}$$

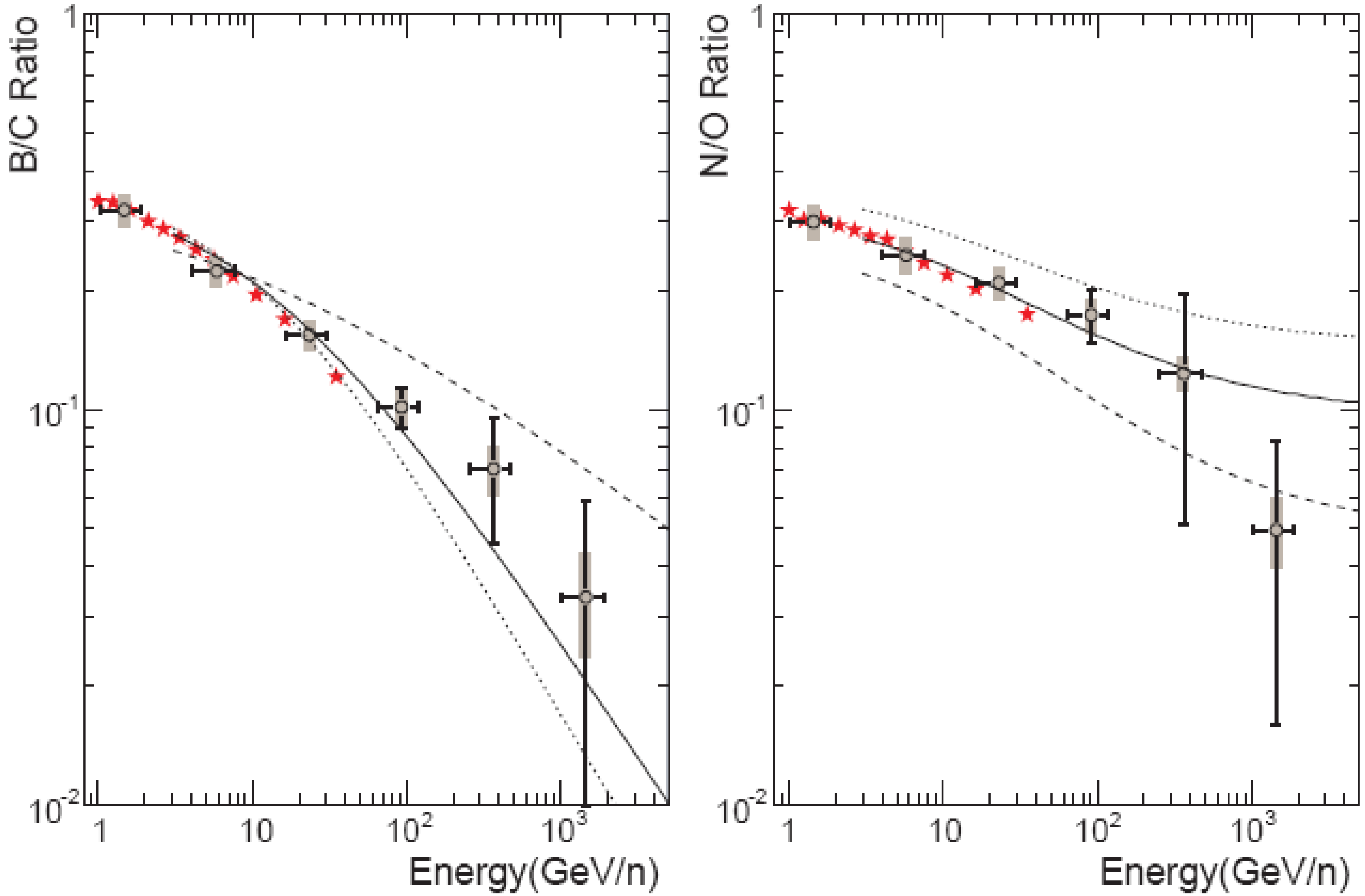
$$\rightarrow \tau_{conf}(E) \propto E^{-x}$$

- Kolmogorov spectrum $\rightarrow \tau_{conf}(E) \propto E^{-1/3}$
  - $x - 2 = 1/3 < 0,7$ … clearly not enough but…
  - ISM perturbations ? Diffusion-convection, MHD ?
- Determination of $\tau_{conf}(E)$ a posteriori :

$$2,7 - 2 = 0,7 \;\; !!!$$

The CREAM results

Secondary/Primary ratio



Compatible with propagation models with escape terms in $E^{0.6}$

CR anisotropy

- Very weak observed anisotropy (first angular harmonic)
  - $\delta \sim (0,5 - 1)10^{-3}$ de 10^{12} à 10^{14}eV
  - $\delta \sim 1$ à 2% à 10^{17}eV
  - Upper limit : 5% à 10^{18}eV, 10% à $10^{18,5}$eV, 30% à 10^{19}eV
  - Constraints on $D(E)$
  - In the diffusive approximation, one can compute d taking into account the position of the solar system wrt the Galaxy.
- Observed anisotropy <u>incompatible</u> with sources distributed like SNR and a diffusion term $D(E) \mu E^{0.6}$ below 10 GeV…

  … but compatible with $D(E) \mu E^{1/3}$!

Sources Distribution

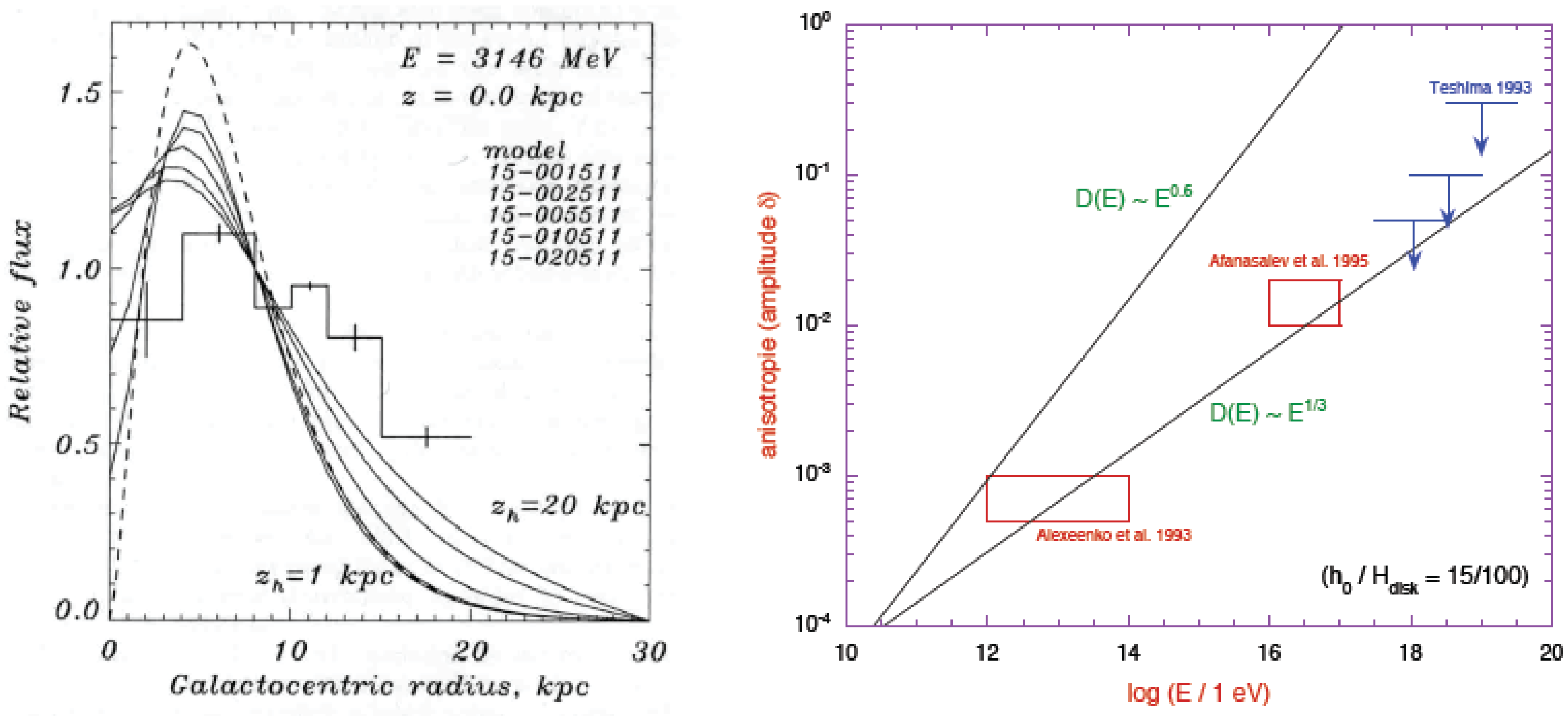


FIG. 5.1 – À gauche : distribution galactocentrique des flux de rayons cosmiques prédits par les modèles de propagations pour différentes tailles du halo de confinement (courbes en trait continu), comparée à la ditribution déduite des observations d'EGRET. La courbe en pointillé montre la distribution source supposée, conforme à la distribution des SNRs, piquée sur l'anneau moléculaire à 4 kpc. À droite : évolution de l'anisotropie du rayonnement cosmique prédite pour un coefficient de diffusion en $E^{0.6}$ (hypothèse SNR) ou en $E^{1/3}$ (comme attendu théoriquement et suggéré par les rapports d'abondance du rayonnement cosmique), comparée aux contraintes observationnelles. La seconde solution est seule en accord avec les données.

® Source distribution is more uniforme...

Full transport equation

$$\frac{\partial \psi(\vec{r},p,t)}{\partial t} = q(\vec{r},p)$$ sources (SNR, nuclear reactions...)

diffusion $+ \vec{\nabla}.\left[D_{\chi\chi}\vec{\nabla}\psi - \vec{V}\psi\right]$ convection

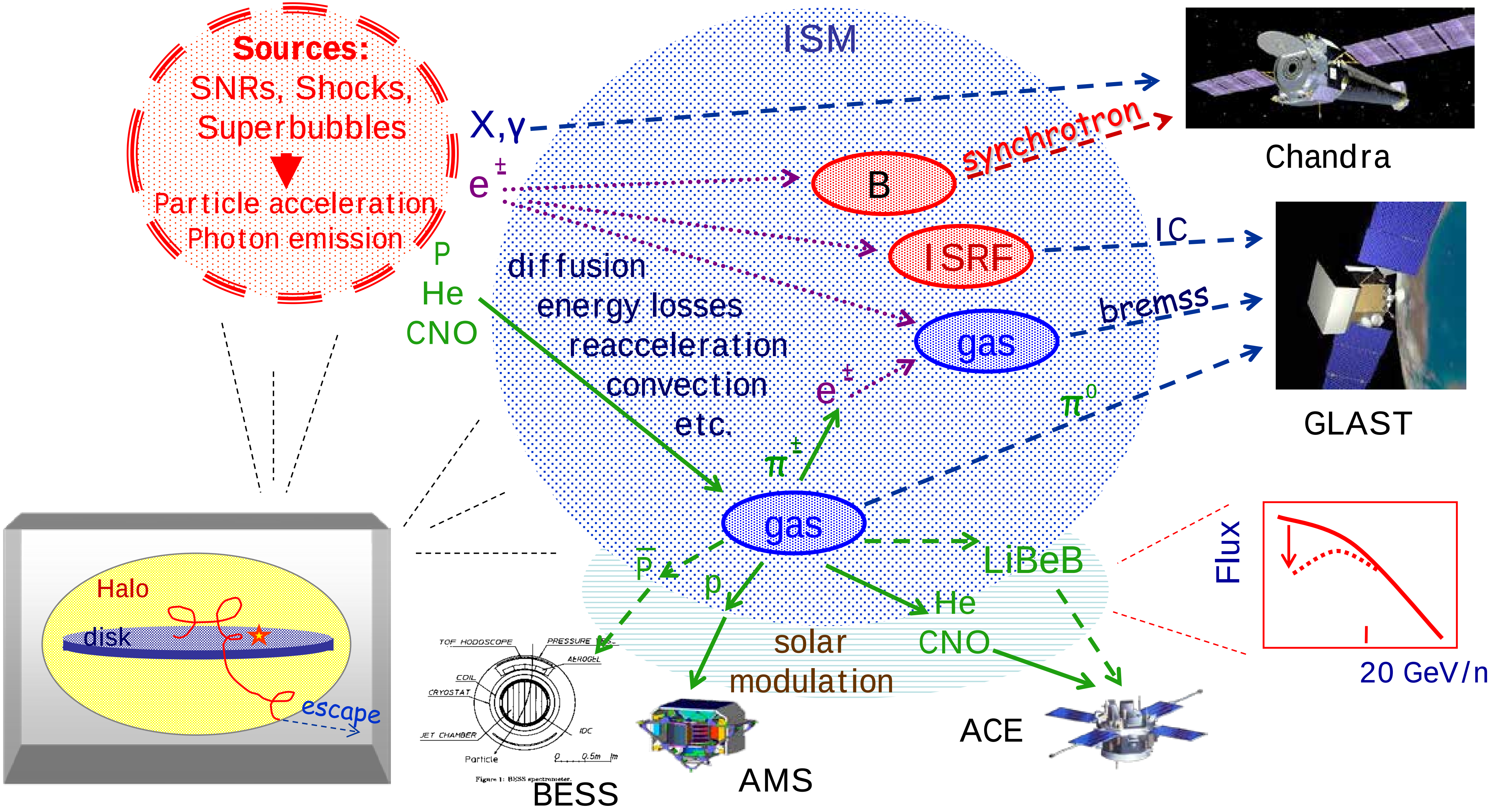
diffusive reacceleration $+ \frac{\partial}{\partial p}\left[p^2 D_{pp}\frac{\partial}{\partial p}\left(\frac{\psi}{p^2}\right)\right]$

E-loss $- \frac{\partial}{\partial p}\left[\frac{dp}{dt}\psi - \frac{1}{3}p\vec{\nabla}.\vec{V}\psi\right]$ convection

fragmentation $- \frac{\psi}{\tau_f} - \frac{\psi}{\tau_d}$ Radioactive decay

*ψ(**r**,p,t) – momentum density*

Propagation in the ISM
et observational contrains



Summary for galactic CR

Everything works fairly well...

- Propagation in the ISM:
  - Complete theory with energy losses, diffusion, in flight nuclear reactions, CR escape, reacceleration, ... impressive results.

(see for example GLAPROP model, A. Strong et I Moskalenko)

- Secondaries / Primaries
- Cosmic clock
- Anisotropies
- Theoretical expectations (~ Kolmogorov spectrum : $D(E)$ µ $E^{0.36}$)

...except naïve acceleration models!

- Observation + models require source spectra / $E^{-2.35}$ (high energy spectral shape and I I^{aires}/ I^{aires} ratio "best fit)
- "Softer" (steeper) than standard spectra for strong shocks f (E^{-2})

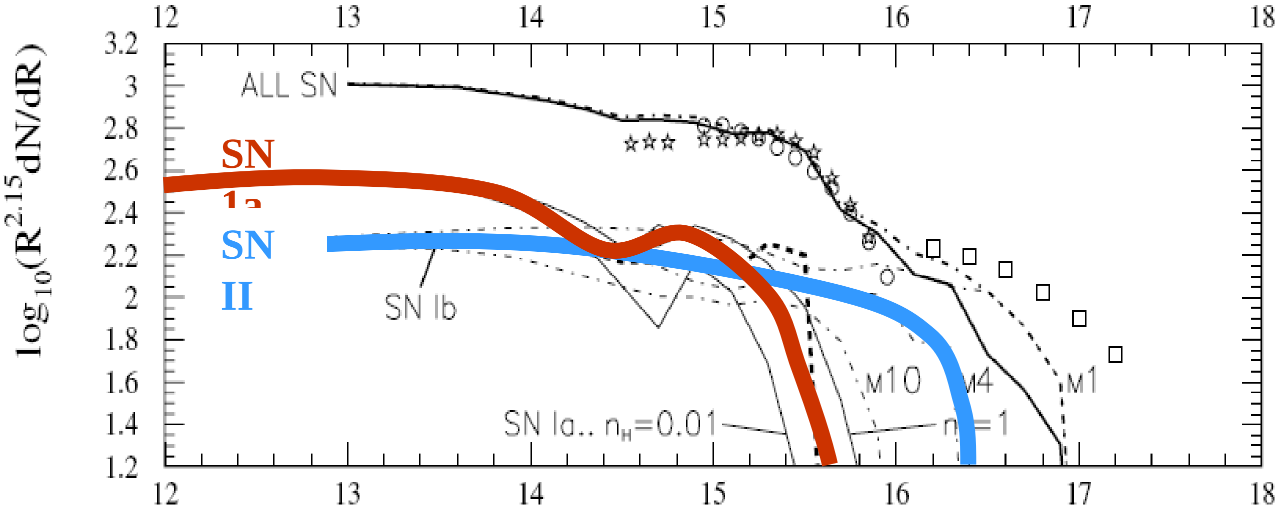
It is possible to find an agreement between diffusive propagation models and standard SNR models,

- Cut off energy, knee, non-linearities, °-ray emission by SNR, source distribution...

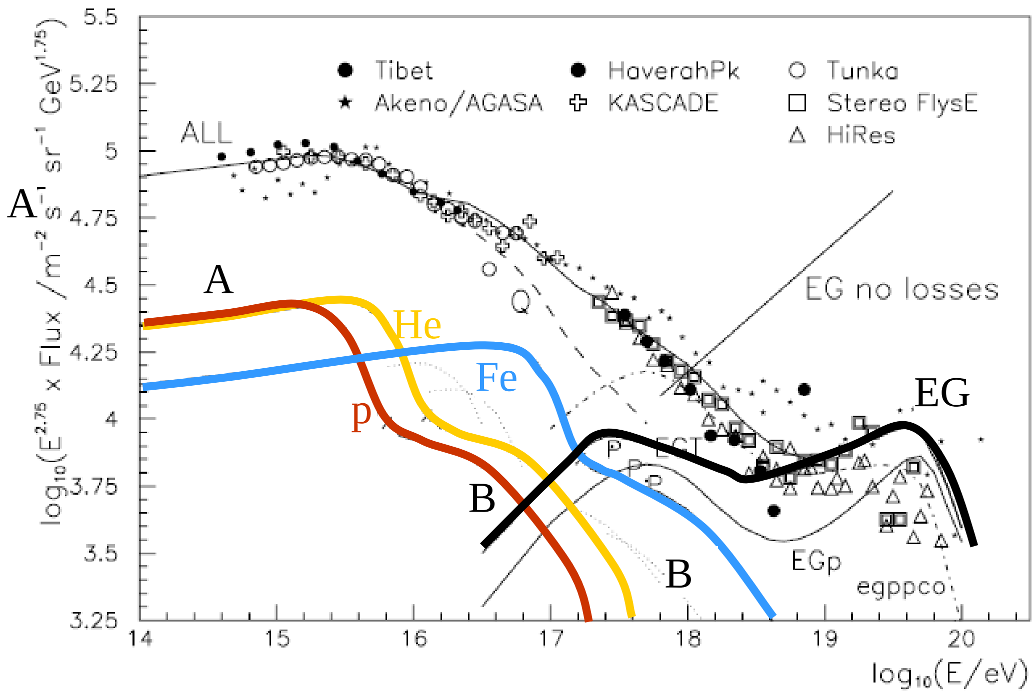
Many parameters) need many observational constrains.

A conciliation exemple CR SM & diffusion

M.Hillas J. Phys., Conf. Ser. 47 (2006) 168 (http://iopscience.iop.org/1742-6596/47/1/021)



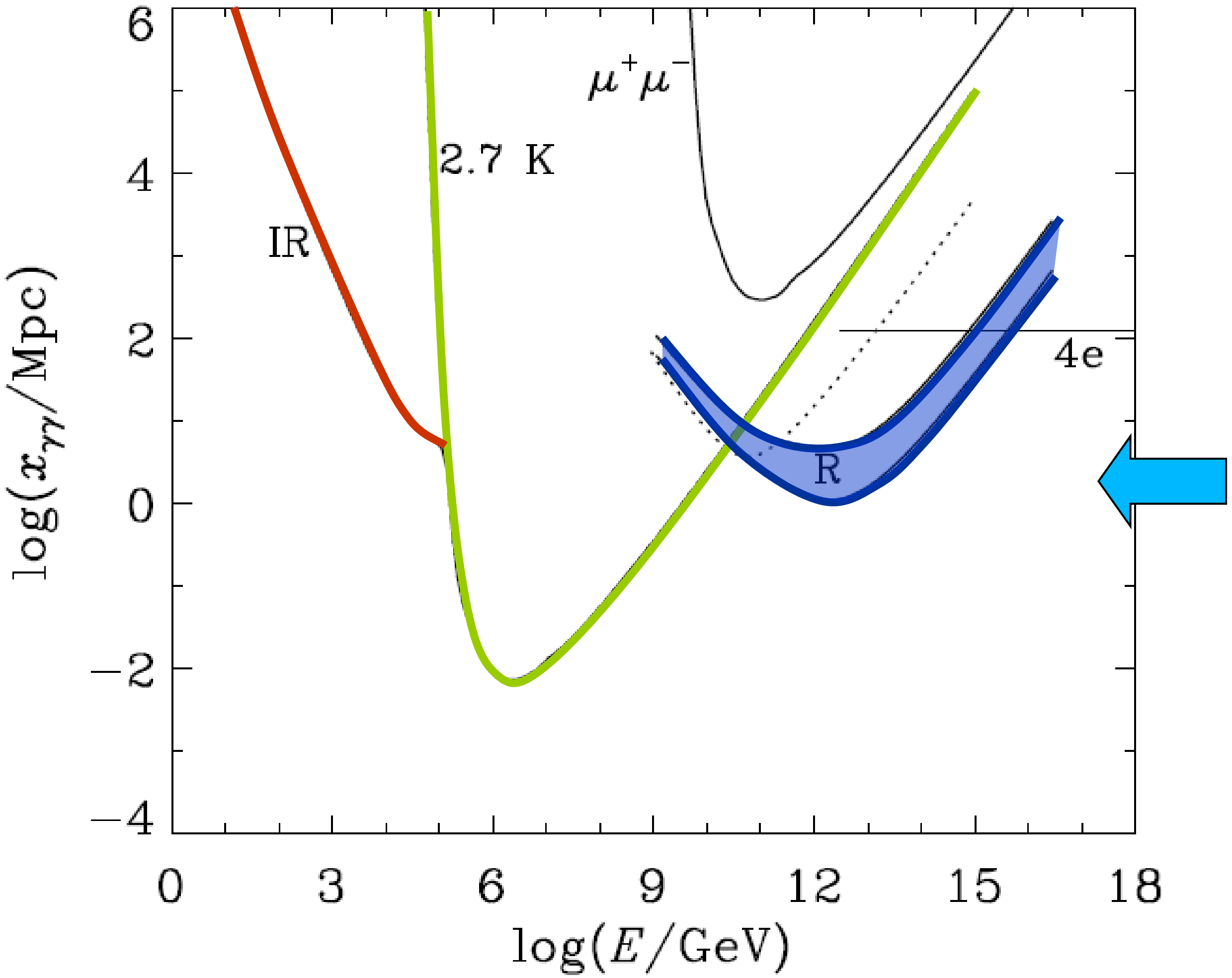
Based on former work by Bell et Lucek on amplification and self-production of Bfield perturbations by CR.



GAMMA-RAY PROPAGATION

Photon atténuation at VHE by intergalactic photon backgrounds

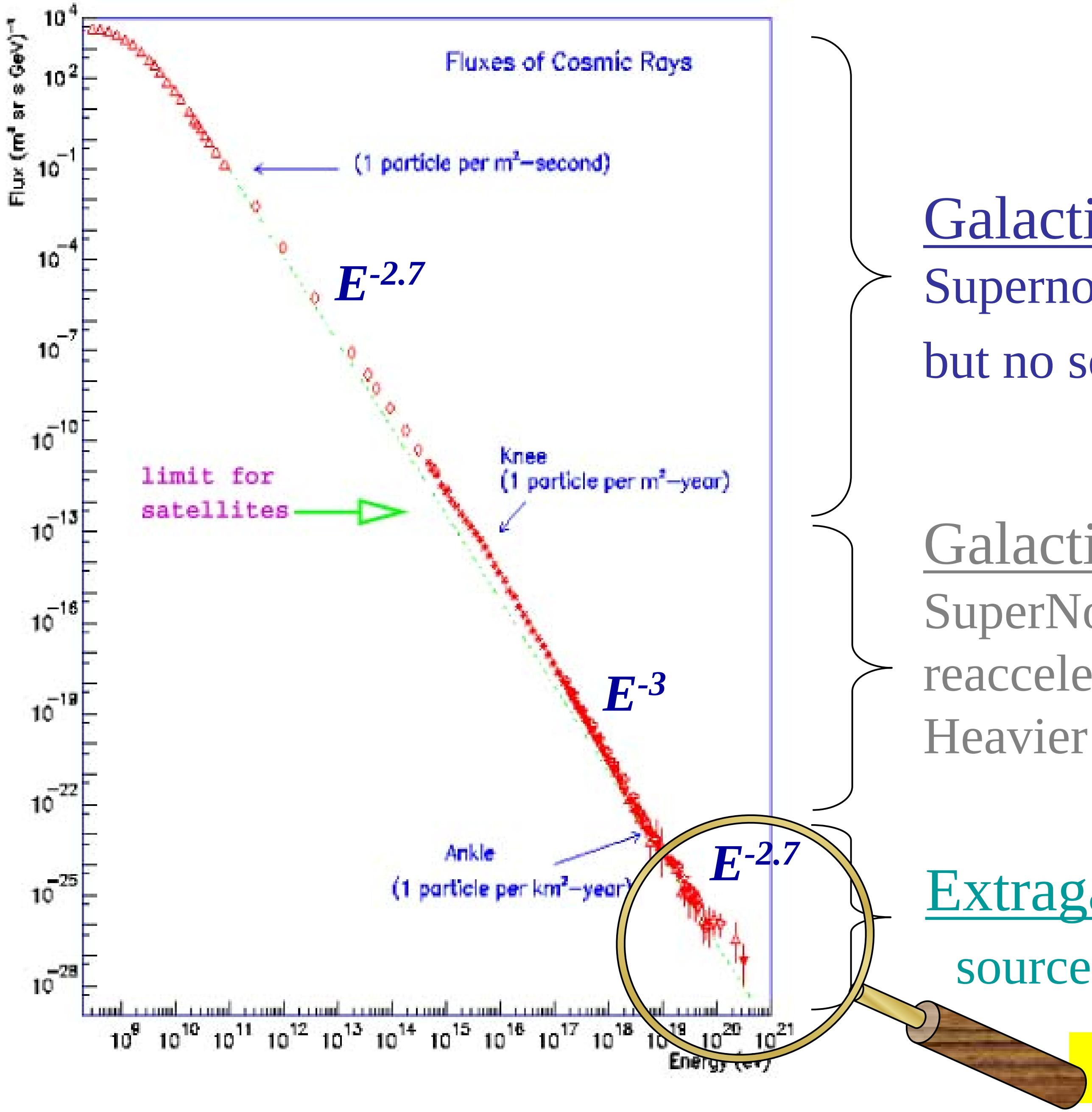
$$\gamma + \gamma_{2.7K} \rightarrow e^{+} + e^{-}$$



Effective γ horizon:
100Mpc at 1TeV
1Mpc at 10 TeV
and above

UHECR PROPAGATION

The CR spectrum



Galactic CR :
Supernovae, MIS,
but no source pointing!

Galactic ?
SuperNovae? Superbubles?
reacceleration?
Heavier nuclei ® protons ?

Extragalactic ?
source ?, composition ?

UHECR, terra incognita

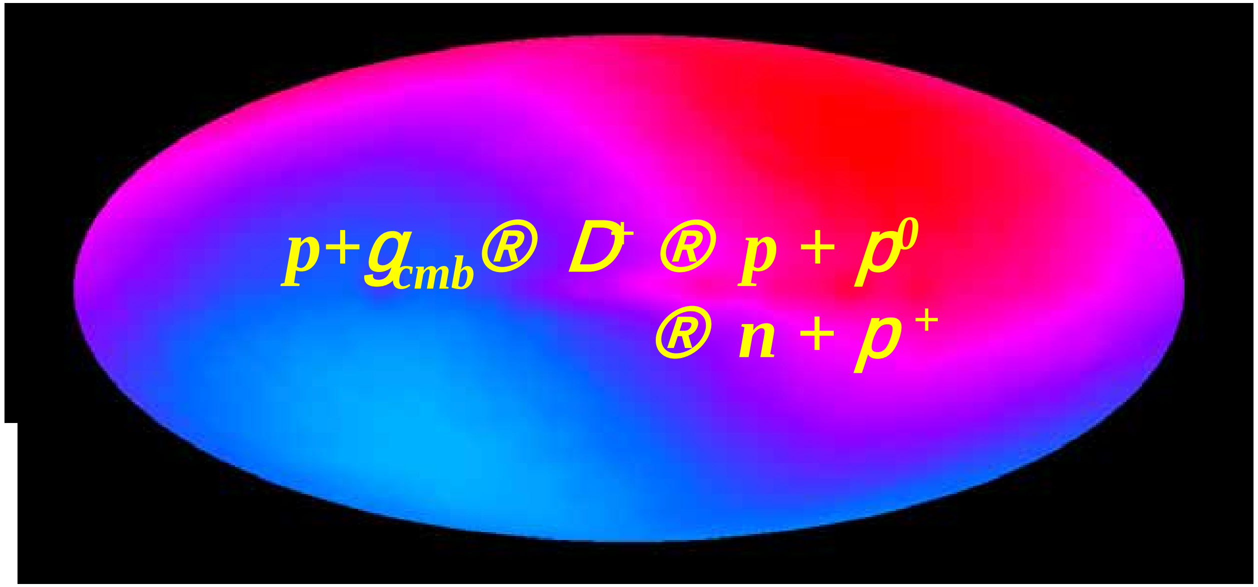
AUGER

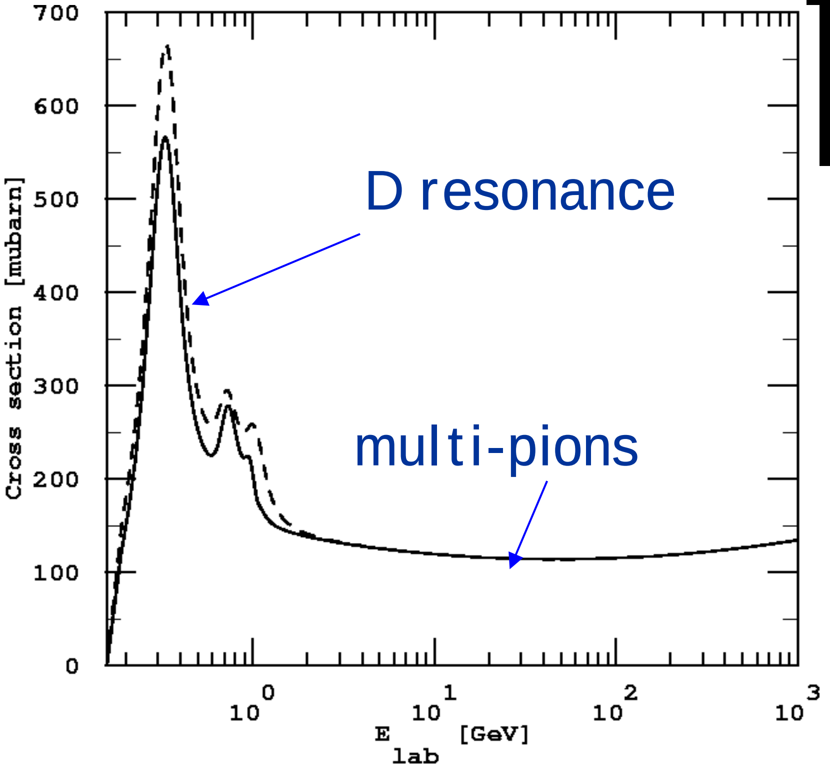
UHECR propagation

3 essential effects :

- Energy losses: modify the spectral shape
- Particle confinement
  (escape depending on energy)
- Spatial and angular diffusion
  due to magnetic fields.
  (regular or fluctuating, inhomogeneities, waves)

An extrem case of relativistic kinematics !!!



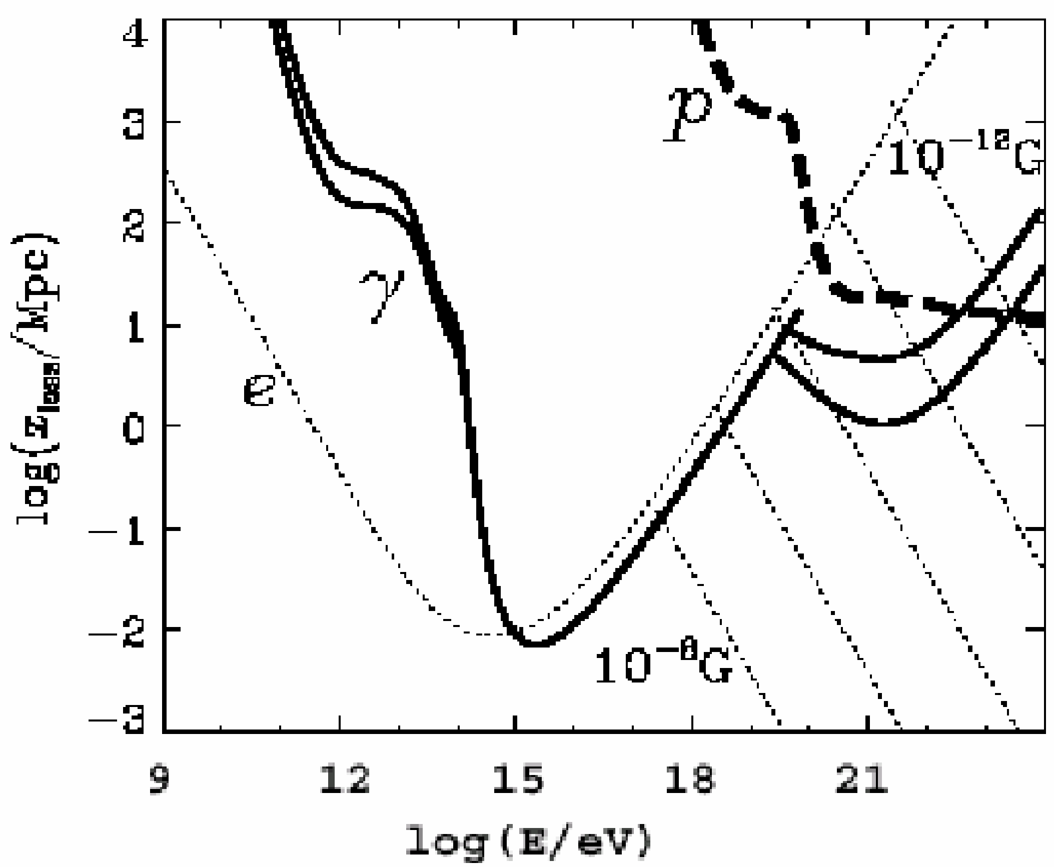


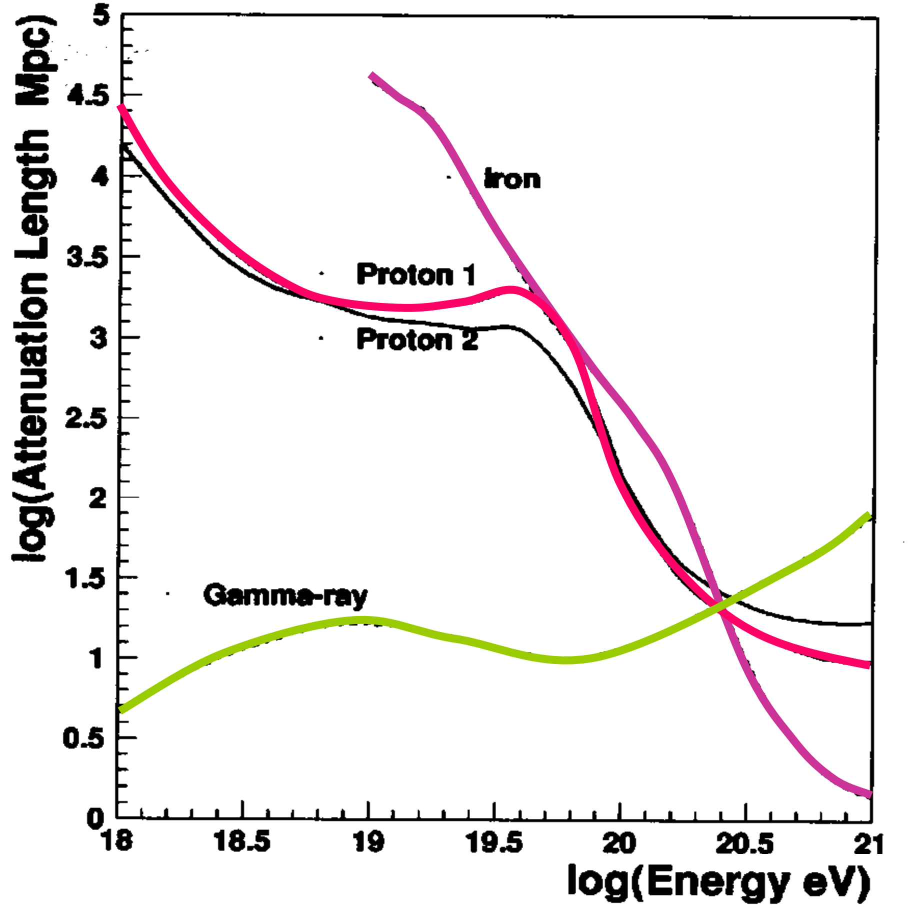
GZK "cutoff"

Greisen '66, Zatsepin & Kuzmin '66

Energy losses

- $p + \gamma_{2.7K} \rightarrow n + \pi^{+}; p + \pi^{0}; p + e^{+} + e^{-}$
- $A + \gamma_{2.7K} \rightarrow (A-1) + N; (A-2) + 2N; A + e^{+} + e^{-}$
- $\gamma + \gamma_{2.7K} \rightarrow e^{+} + e^{-}$





Consequences on spectral shape

Protons:

Photo Pion production
CMB photons

$p+g_{cmb}$® D® p/n+p

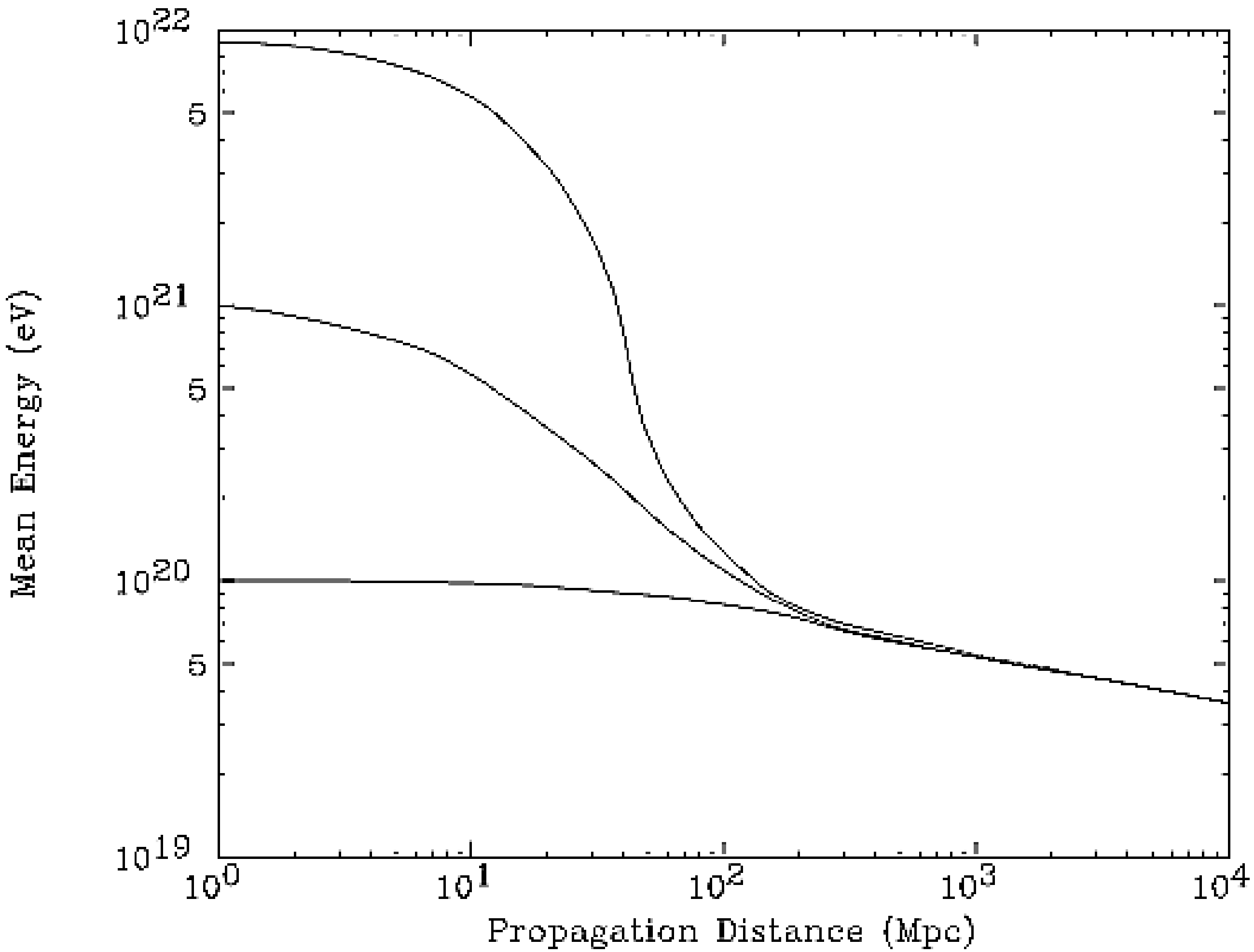
Fe:

photo-dissociation

on IR bg and CMB

Photons: paires $e^{+}e^{-}$

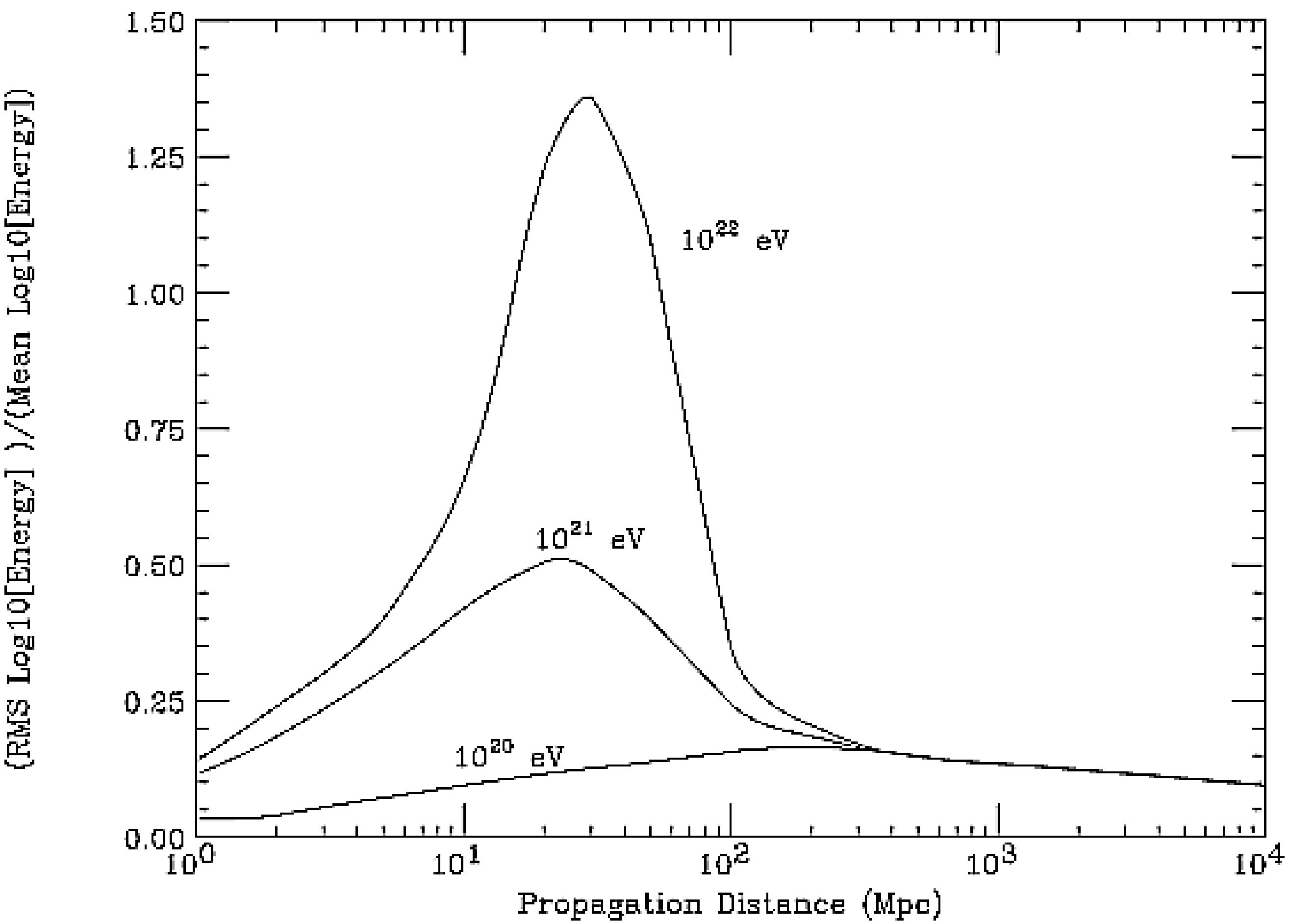
on radio bg



Protons energy vs. distance (J. Cronin)
Energy loss on CMB

UHE Extragalactic Particles

Fluctuation dues to multiple scattering



Protons energy vs. distance (J. Cronin)
Energy loss on CMB

UHE Extragalactic Particles

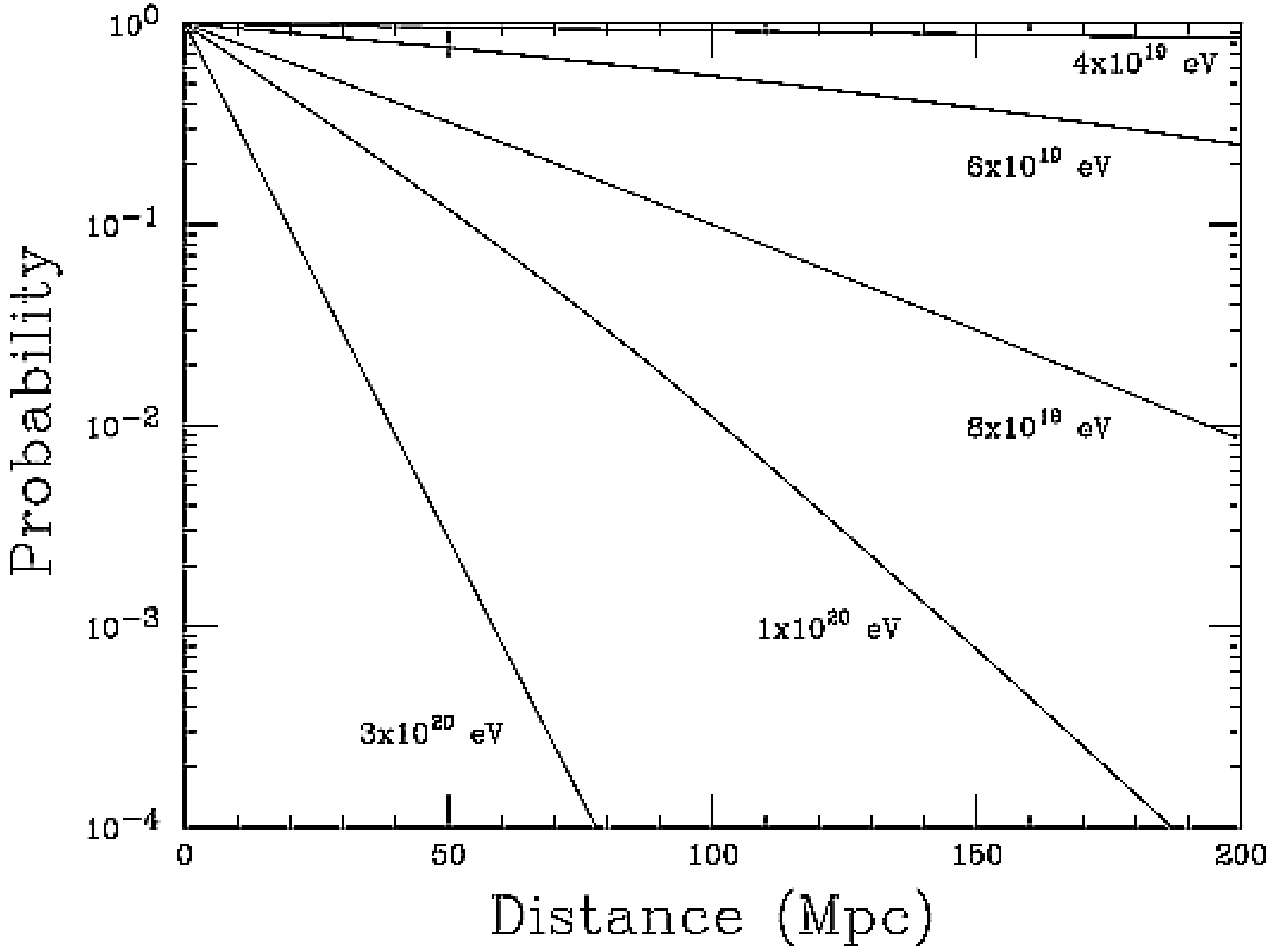
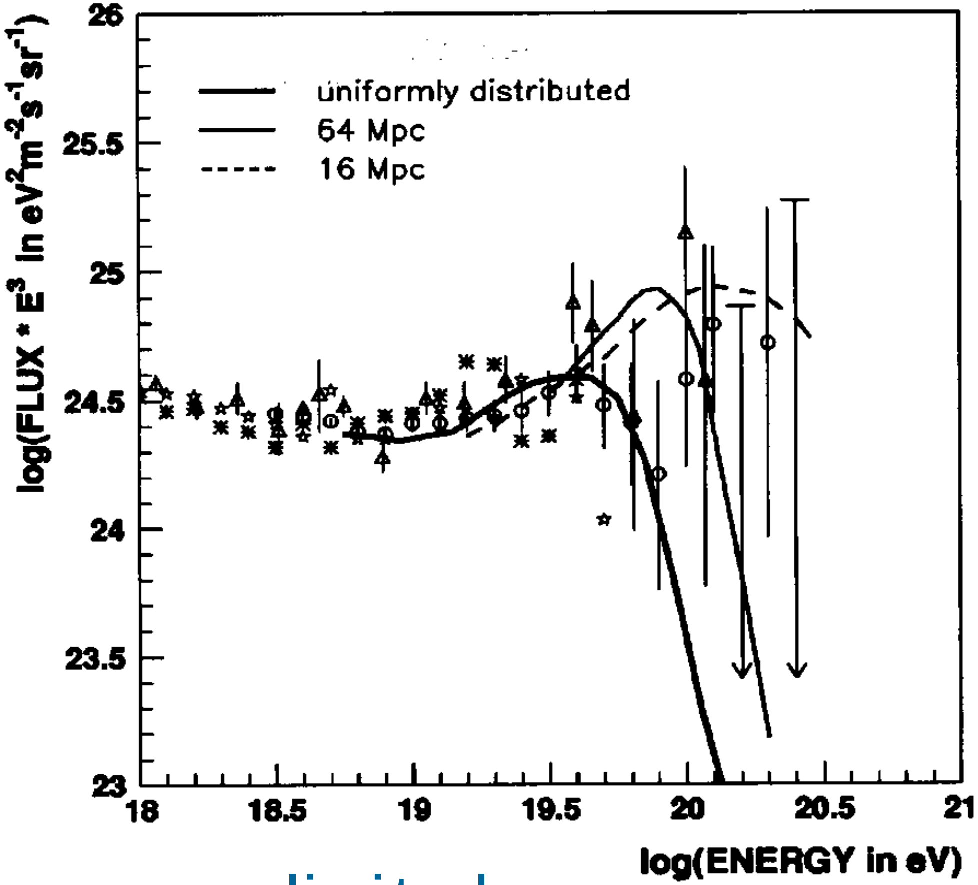


Figure 6: Probability that an observed event at a given energy has its source at a distance greater than the indicated distance. A source spectrum proportional to $E^{-2.5}$ is assumed. Figure provided Paul Sommers, University of Utah.

Coupure GZK

- Greisen-Zatsepin-Kuz'min



- Distance to the source limited to 100 Mpc for 10^{20} eV protons, and 15 Mpc for 3´ 10^{20} !
- Actually even worse if particles are deflected ($D_{effectif} > D_{linear}$)

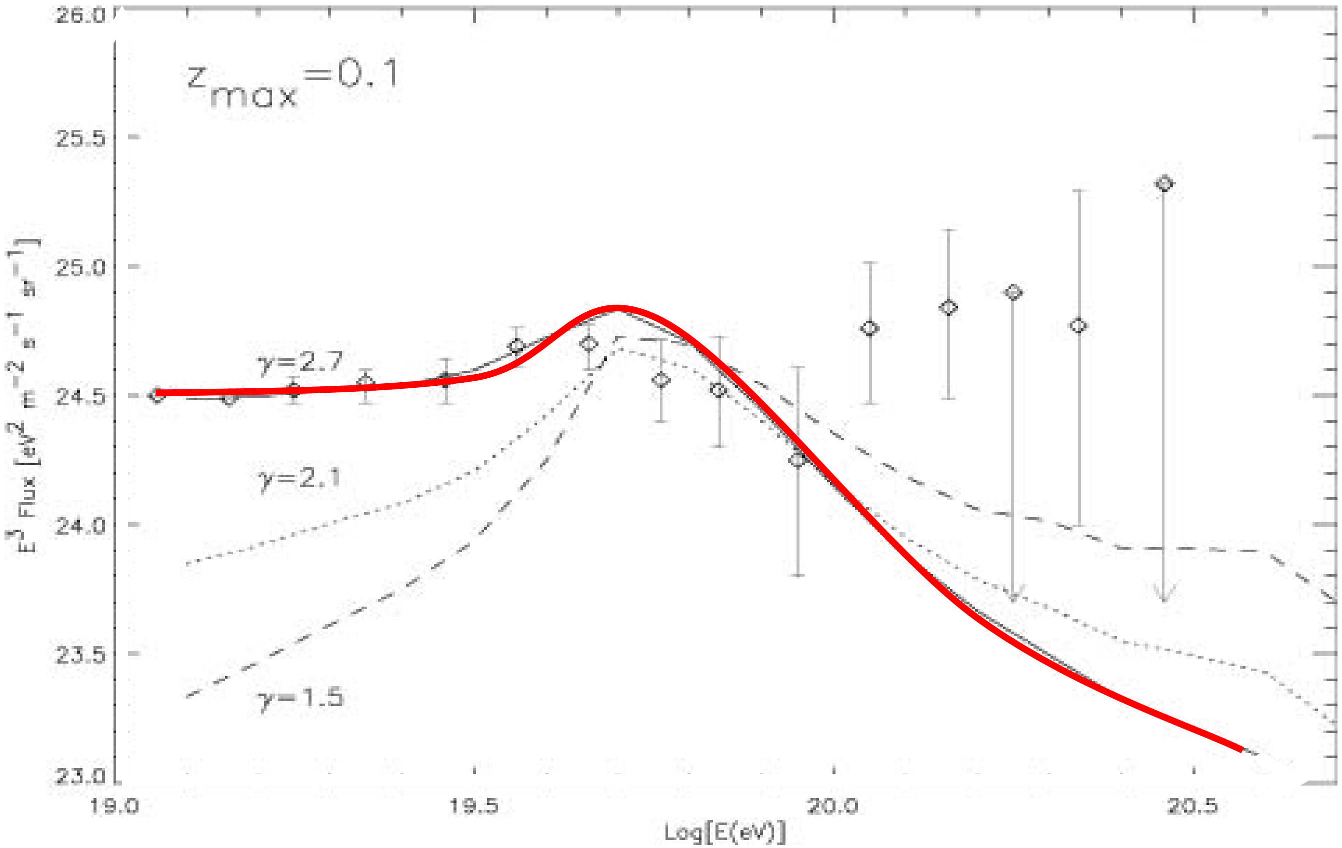
Spectrum above GZK cutoff

Depends among other things:

- From the injection spectrum
- From the sources cosmological distribution (evolution)
- From IG magnetic fields

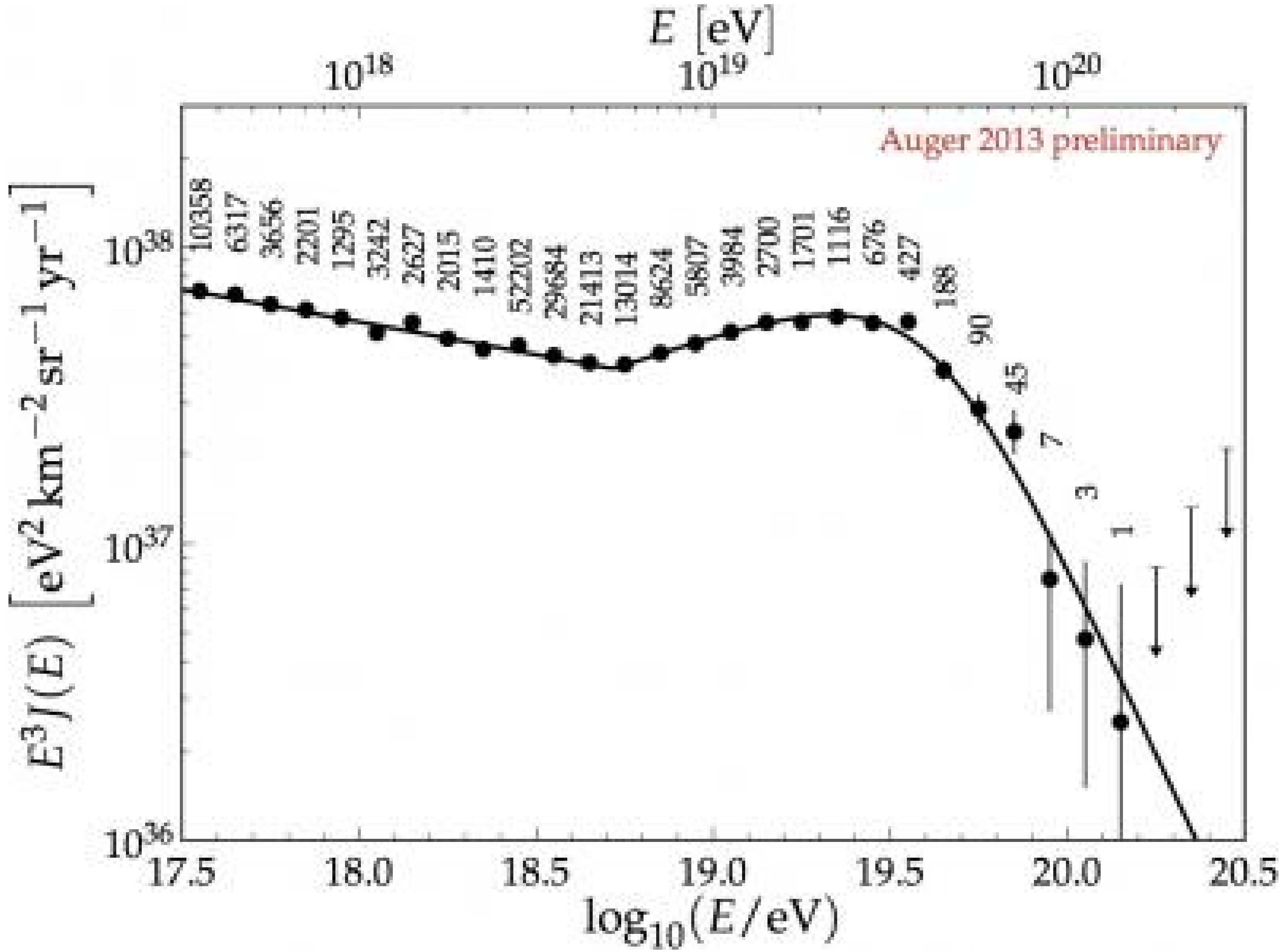
-

E^3 Flux

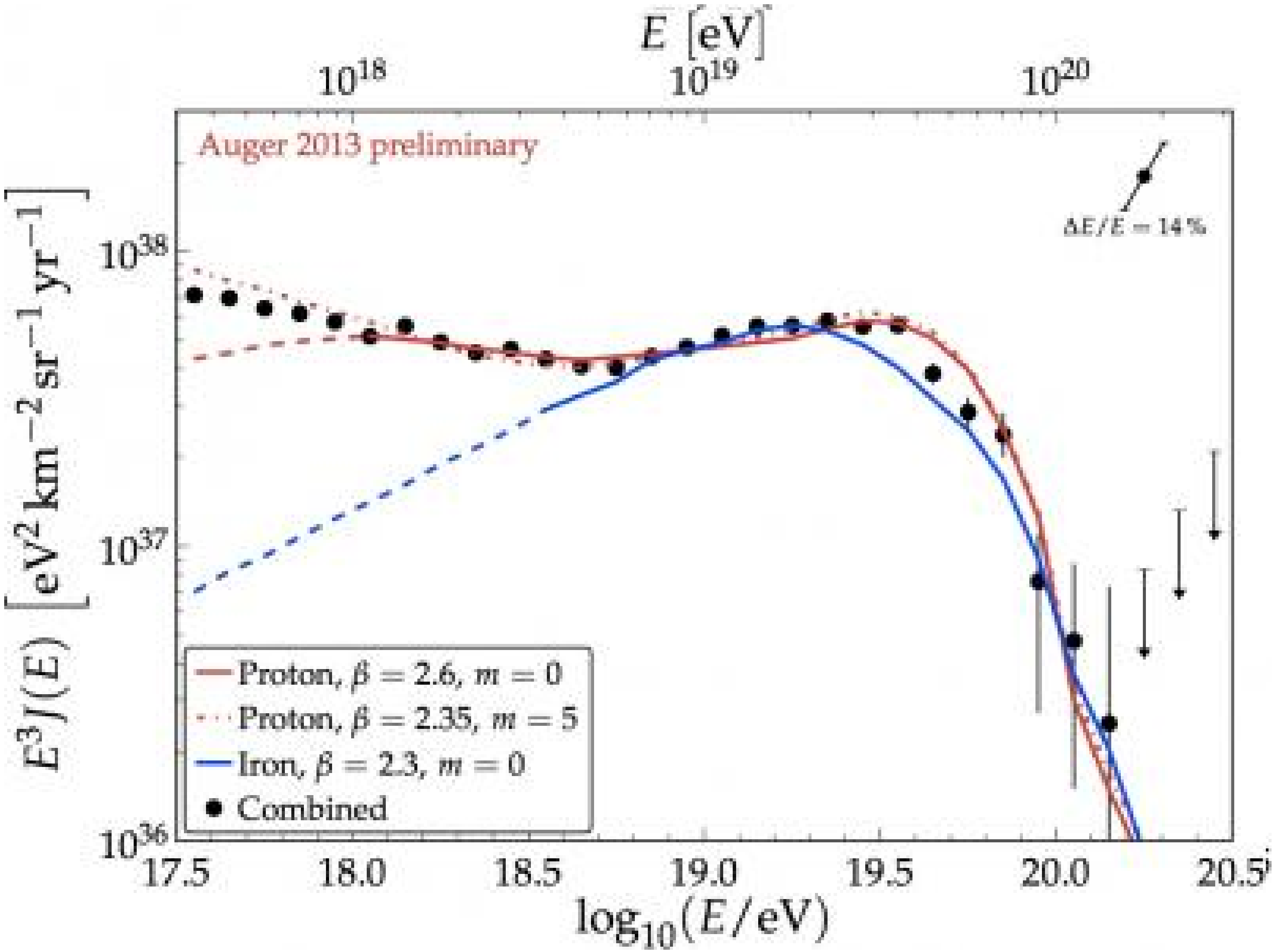


Blasi, Olinto. '01

GZK like suppression (Auger)



GZK like suppression (Auger)

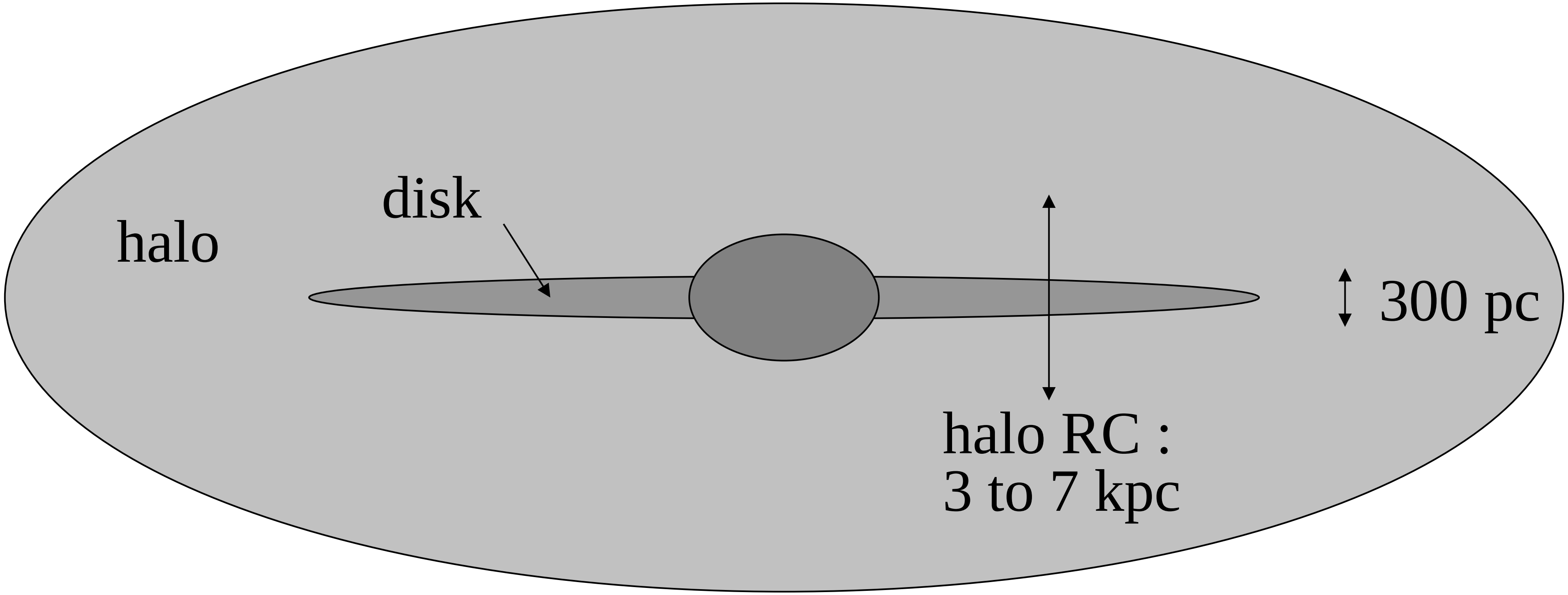


F.Montanet Experimental Astroparticle Physics ESIPAP 2014

MAGNETIC DEFLECTIONS

Galactic magnetic deflection

- 10^{18} eV proton in a B = 3 μG field Þ r_g ~ 370 pc

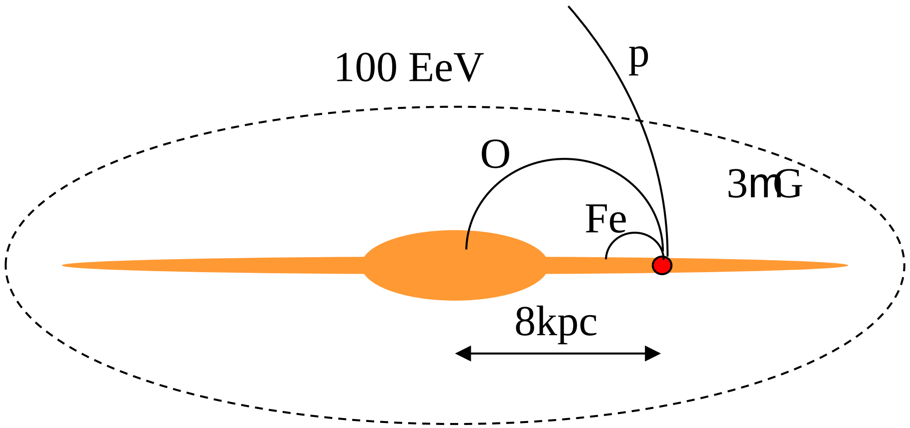


- 2´ 10^{19} eV proton in B = 3 μG Þ r_g ~ 7 kpc
- 5´ 10^{20} eV Fe in B = 3 μG Þ r_g ~ 7 kpc

Propagation the Galaxy

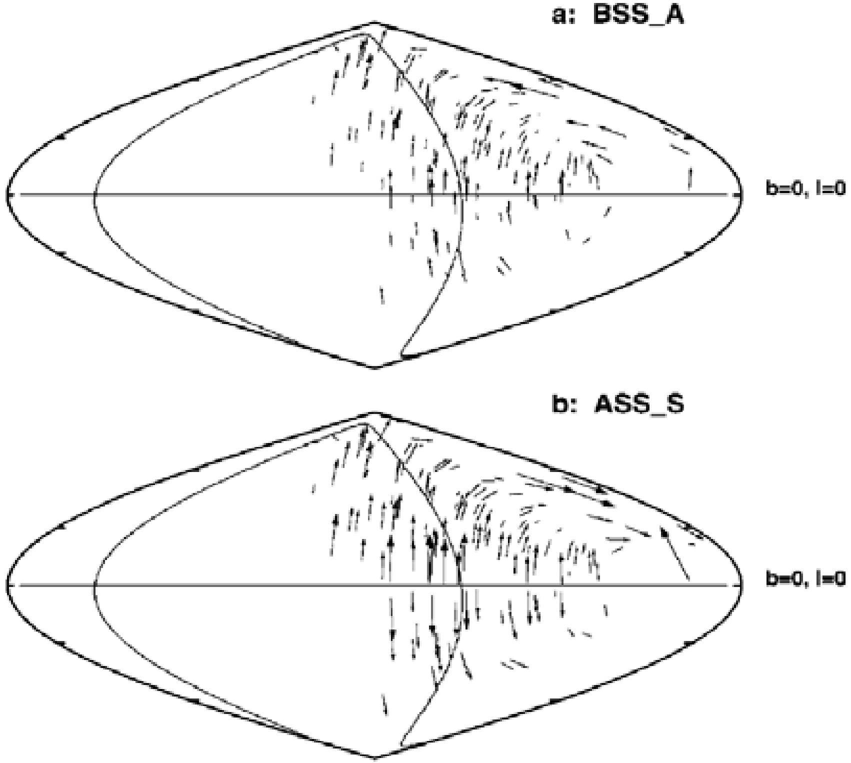
- Galactic magnetic field model

$$\left(\frac{R_{Larmor}}{kpc}\right) = \left(\frac{1}{Z}\right) \times \left(\frac{E}{1EeV}\right) \times \left(\frac{ZB}{\mu G}\right)$$



- Possible galactic confinement of 10^{20}eV nuclei

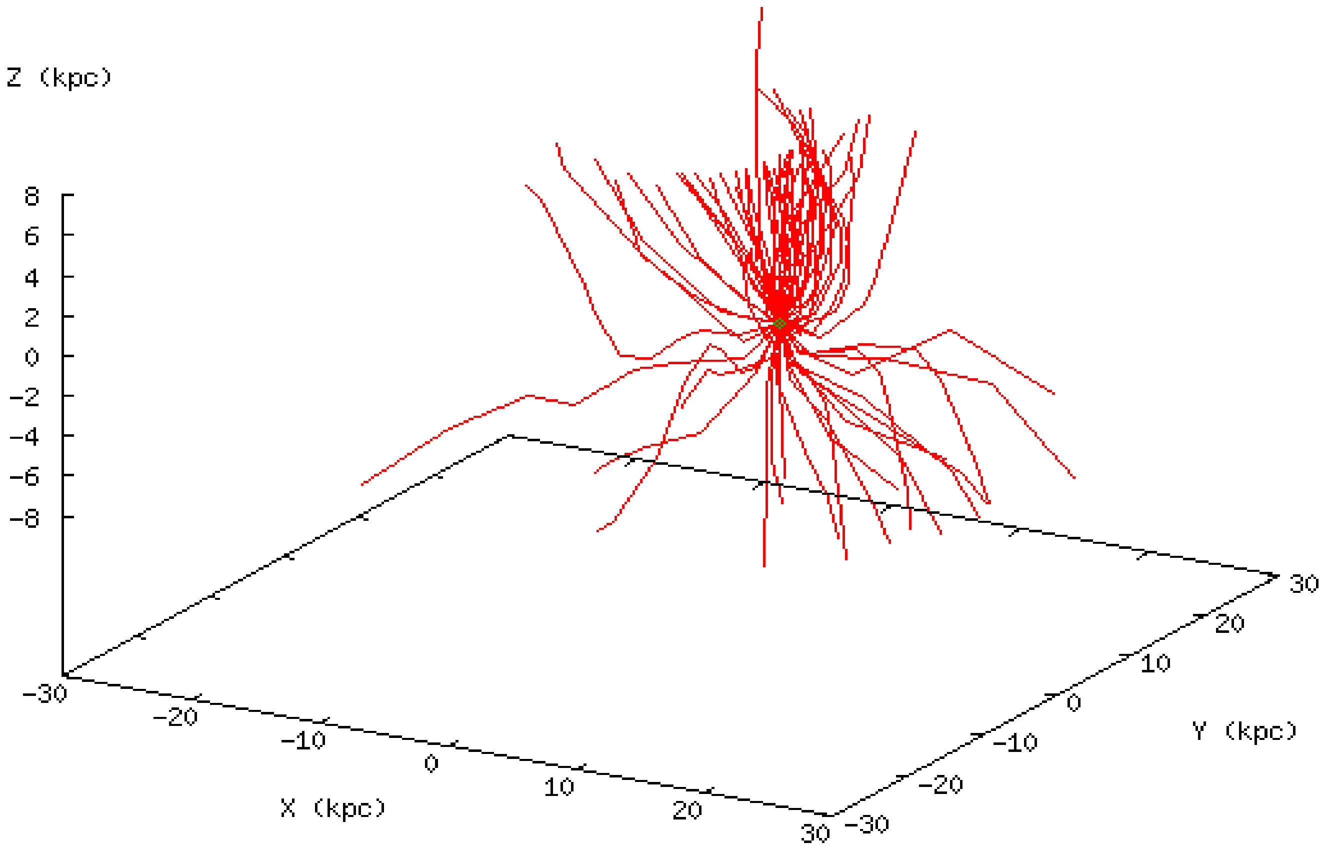
- 10^{18}eV neutrons decay length
  $\beta\gamma c\tau \approx 10$kpc $\Rightarrow$ galactic distances



Tracking back direction of proton events >4 10^{19} out of the Galaxy, two different field hypothesis [Stanev97]

Pointing at UHECR sources?

100 EeV Iron Nucleus Distribution Under the Influence of Regular Galactic Field and Galactic Wind Field



O'Neil, Olinto, Blasi '01

Pointing at UHECR sources?

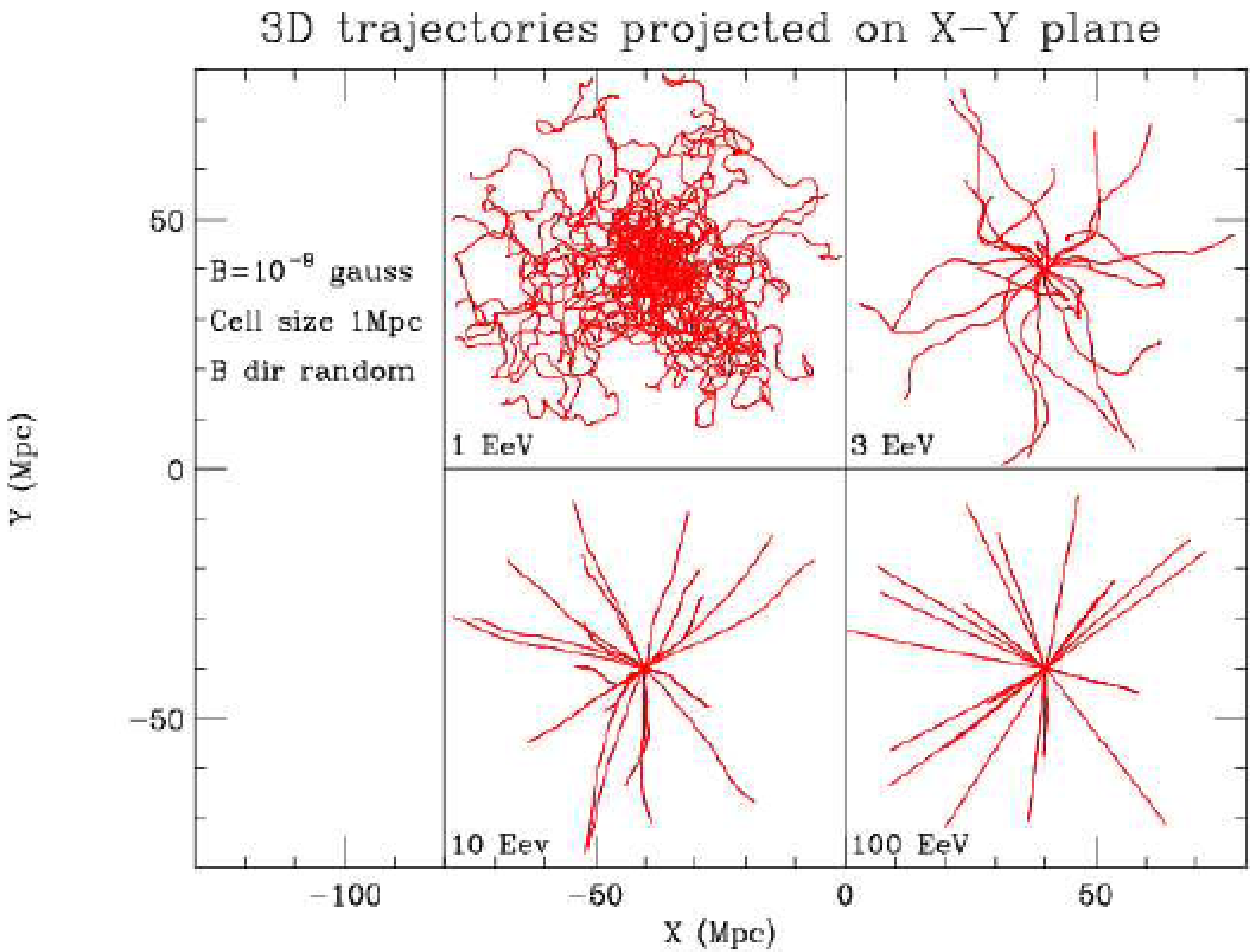


Figure 7: Projected view of 20 trajectories of proton primaries emanating from a point source for several energies. Trajectories are plotted until they reach a physical distance from the source of 40Mpc. See text for details.

Mapping IG fields with UHECR?

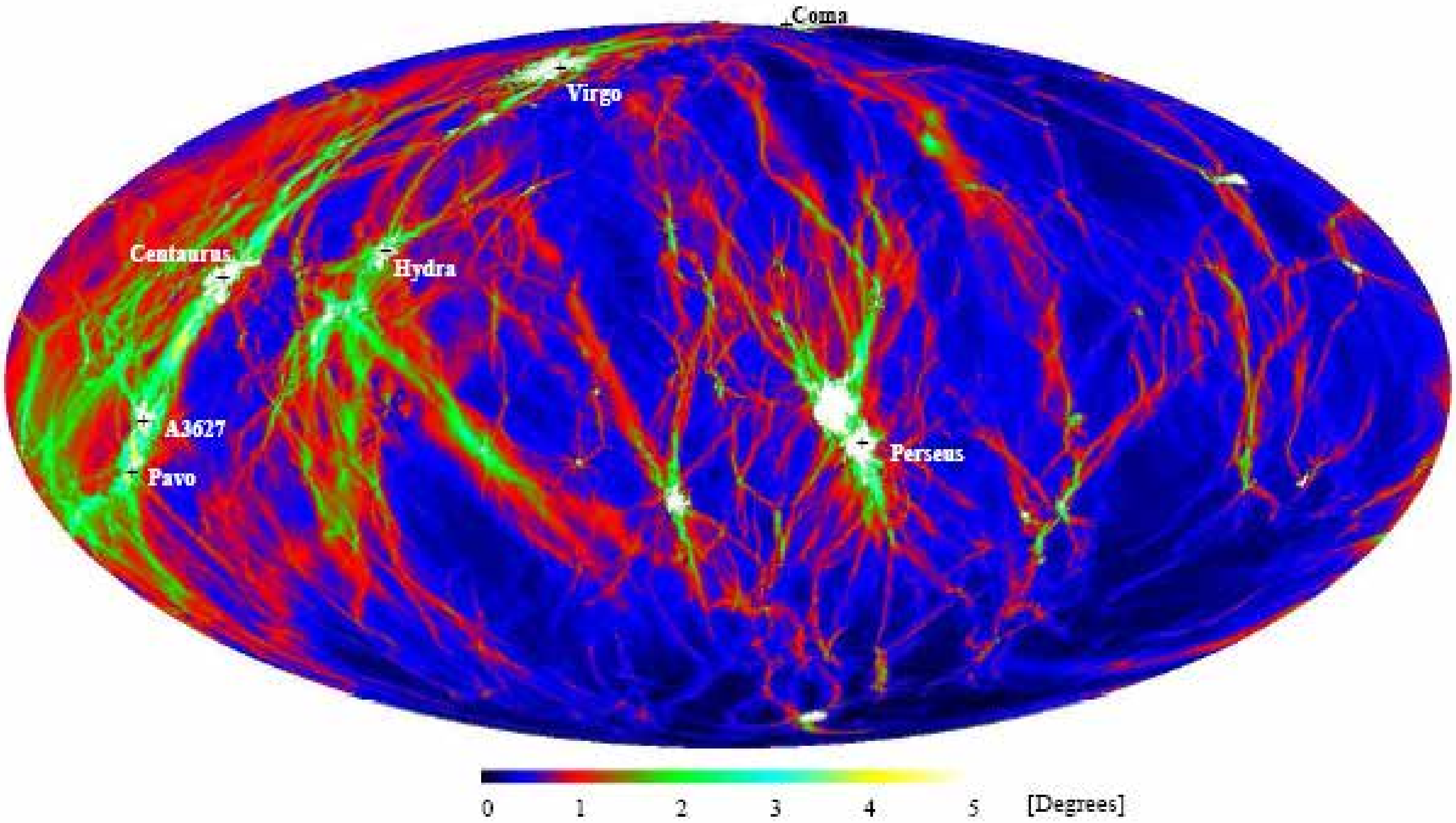


FIG. 1. Full sky map (area preserving projection) of deflection angles for UHECRs with energy 4×10^{19} eV using a linear color scale. All structure within a radius of 107 Mpc around the position of the Galaxy was used. The coordinate system is galactic, with the galactic anti-center in the middle of the map. Positions of identified clusters are marked using the locations of the corresponding halos in the simulation. Note that deflections internal to the Milky Way have not been included.

Diffusion in the Universe

- If they are protons,
  arrival direction ≈ source direction

  $r_L \approx 100 kpc \times Z \times (E/10^{20} eV) \times (B/10^{-6} G)^{-1}$

  $\delta\theta \sim \lambda_B / r_L$ deviation per field correlation length $\Rightarrow \Delta\theta \sim \sqrt{D\lambda_B}/r_L$

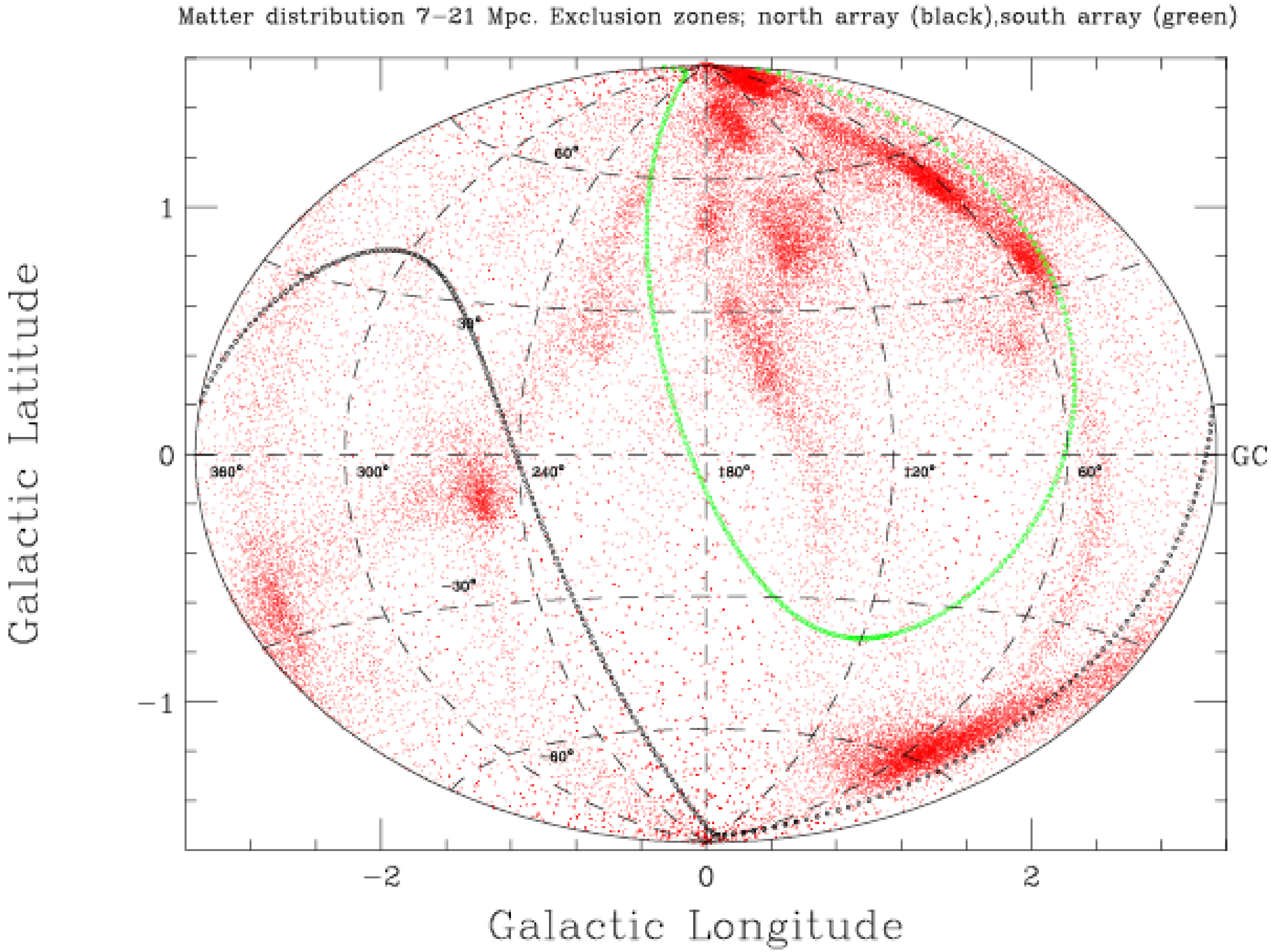
  ® Proton astronomy !

- Correlations between arrival directions and sources:
  UHECR distribution is NOT ISOTROPIC!!! (AUGER 2007)
  - Confirmation of a GZK limited horizon
  - Few sources in the GZK sphere ® anisotropie
  - Astrophysical origine is confirmed!

- Arrival time delay
  - $\Delta t \sim \Delta\theta^2 d/c \sim D^2 \lambda_B / r_L^2 c$
  - If eruptive or transient sources (GRBs, TDs), the must overlap in time (otherwise E(t)!)
  - Multiplets of events from same direction observed but no significant ordering in E or deviation.

® Correlation must be confirmed! (statistics...)

Matter distribution in the GZK sphere



Observables & Observations

PRIMARY RC DETECTION (ON TOP OF ATMOSPHERE)

How to characterize the primary particle?

- Mass m
- Electric charge Ze
- Velocity $v = \beta c$
- Lorentz Facteur $\gamma = E/mc^2$
- Momentum $p = mc\beta\gamma$
- Kinetic energy $T = mc^2(\gamma - 1)$

How to characterize the primary particle?

Detector	Observable	Link with the particule
Magnetic spectrometer	Rigidity & Sign of Z	pc/Ze
Time of flight	Velocity/c	β
Proportionnal counters Scintillators Ionisation chamber	Ionisation	$dE/dx = Z^2 f(\beta)$
Čerenkov effect	Č photons density	$dN/dx = Z^2 g(\beta)$
Transition radiation	Nomber of photons X	$N = Z^2 h(\gamma)$
Calorimeter	Deposited energie	$mc^2(\gamma - 1)$

Two important radiations for particle identification

Two effects of the polarization induced by charged particles in dielectric medium

Proportionnal to Z^2

- Čerenkov radiation : si $v > c/n$
  Sensitive to $\beta = v/c$
- Transition radiation : at the interface of $\neq$ dielectric media
  Sensitive to $\gamma = E/(mc^2)$

Cherenkov Radiation

- Emitted on a cone whose axis is along the particle trajectory and of half oppening angle θ_C tel que $\cos\theta_C = 1/(\beta n(\omega))$
- Threshold defined by the condition $\cos\theta_C < 1$
- Emission at all frequencies with $n(\omega) > 1$ (from UV to radio), flat in $h\nu$ in photon yield Generaly detected from near UV to visible light.

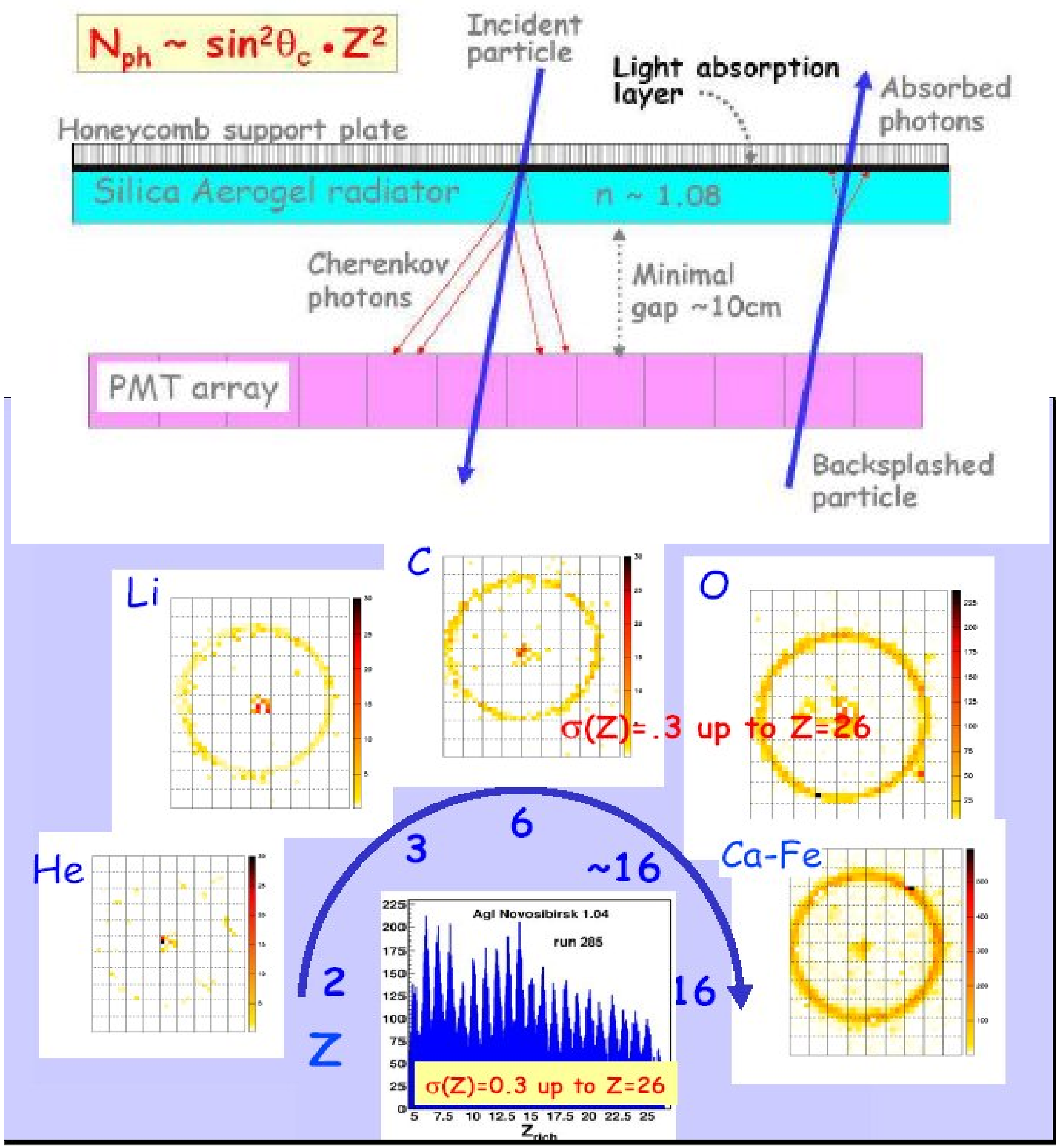
$$\frac{dN^2}{dx d\omega} = \frac{\alpha}{c} Z^2 \sin^2\theta_C$$

- Allow discrimination between particles with same momentum and $\neq$ masses. (électrons / protons /noyaux) up to energies of the order of 10 GeV/nucleon if $\Delta\beta/\beta \approx 10^{-3}$
- Allow determining the orientation of the paricle direction.
- "Ring Imaging Cherenkov" detector or RICH : allow a precise mesurement of the charge.

Cherenkov imaging (RICH) and charge measurement

RICH principle →

AMS 2 Prototype →



Transition radiation

- Origine : if a particle traverse the boundary between 2 $\neq$ dieletric media, the solution corresponding to each medium do not satisty the conditions at the boundary.
  $\rightarrow$ nécessité d'une onde «libre» supplémentaire.

- A dielectric medium is caracterised by its "plasma" frequency ω_p (oscillation frequency of free-like electrons)
  $\omega_p = \sqrt{\frac{n_e e^2}{\varepsilon_0 m_e}}$ or $\hbar\omega_p = 2E_R\sqrt{4\pi n_e a_B^3} \approx 20\text{eV}$
  $n_e =$ electron density ;
  $E_R =$ Rydberg energy $= 13.6\text{eV}$;
  $a_B =$ Bohr radius

- Roughly half of the energy is emitted in the frequency domaine $0.1\gamma\omega_p < \omega < \gamma\omega_p$ for a Lorentz factor $\gamma \approx 1000$, This is the X-ray domaine (2 to 20 keV)

- Energy emitted by the interface if $I = \alpha z^2 \gamma \hbar\omega_p / 3$

Transition radiation (cont)

- Angular distribution peaked at small angles around particle direction: $\theta \approx 1/\gamma$
- Small yield of X-photons per interface : $N \approx \alpha Z^2 \approx 10^{-2} Z^2$
  $\rightarrow$ multiply the number of interfaces
  $\rightarrow$ stack plastic sheets or fibers
- X-ray detection by photo-electric effect: proportional tubes
- Discriminates between particles with same energy and $\neq$ masses at high energies (100 GeV to 1 TeV) (instrumental detection thereshold for X-rays).
- Can measure the Lorentz factor γ up to 10^5. In this case, choose material adequately to have a progressive threshold.

Charge particles and cosmic antimatter

Satellite born experiment : AMS-1, PAMELA, AMS-2

- First satellite based experiments on cosmic rays : HEAO-C, Ariel-VI (1979) → relatively low energy (up to a few 10 GeV/nucleon)

- First satellite based magnetic spectrometer en satellite :
  AMS-1 on the space shuttle « Discovery » (1998)

- New generation of experiments:
  PAMELA (since Juin 2006)
  AMS-2 (since May 2011)
  on the International Space Station → data up to ~TeV
  and precise measurements of flux of cosmic antiparticles

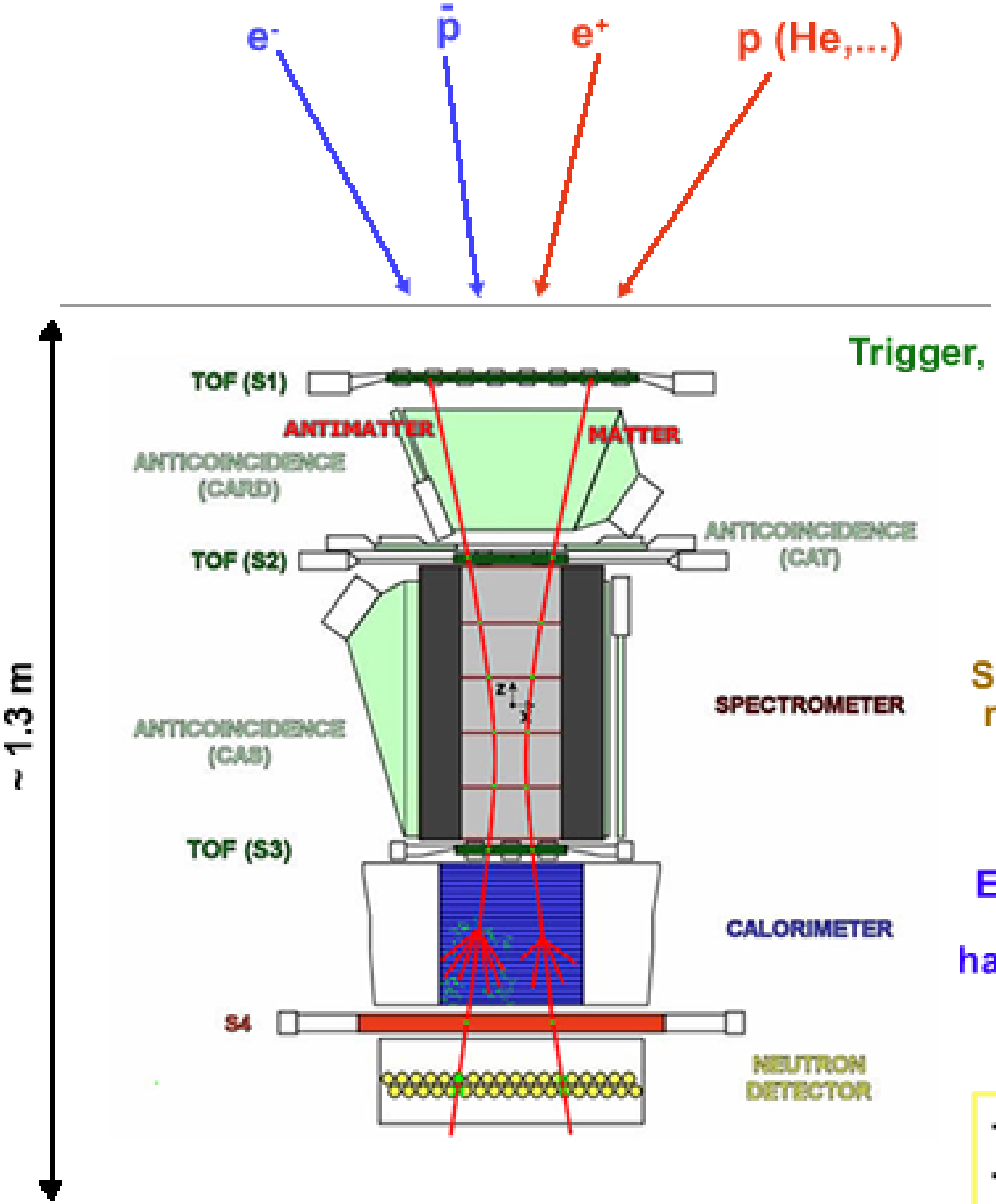
AMS-2 On Board ISS

Mission Number: STS-134
Launch: **May 19, 2011**
Orbiter: Endeavour

Space spectrometers

	AMS-1 (June 1998)	PAMELA (June 2006 - ...)	AMS-2 (May 2011 - ...)
Spectrometer Acceptance	0.82 m^2 sr	20.5 cm^2 sr	0.82 m^2 sr
Spectrometer	Aimant permanent Nd Fe B 0.15 T BL^2 = 0,15 T m^2 6 plans (Si)	Aimant permanent Nd Fe B 0.48 T BL^2 = 0,10 T m^2 6 plans (Si)	Aimant permanent Nd Fe B 0.15 T BL^2 = 0,15 T m^2 6 plans (Si)
Time of Flight	yes	yes	yes
Cherenkov	Aerogel (threshold)	-	Ring Imaging Ch.
Transition rad	-	yes	yes
Neutrons det.	-	^{3}He	-
Anticoincidence	-	yes	yes
Calorimeter	-	16,3 X_0 W+22 plans (Si)	16 X_0 Pb+fibers sc.

PAMELA



~470 Kg / ~360 W

Trigger, ToF, dE/dx

- S1, S2, S3; double layers, x-y
- plastic scintillator (8mm)
- ToF resolution ~300 ps (S1-3 ToF >3 ns)
- lepton-hadron separation < 1 GeV/c
- S1.S2.S3 (low rate) / S2.S3 (high rate)

Sign of charge, rigidity, dE/dx

- Permanent magnet, 0.43 T
- 21.5 cm^2 sr
- 6 planes double-sided silicon strip detectors (300 μm)
- 3 μm resolution in bending view → MDR ~800 GV (6 plane) ~500 GV (5 plane)

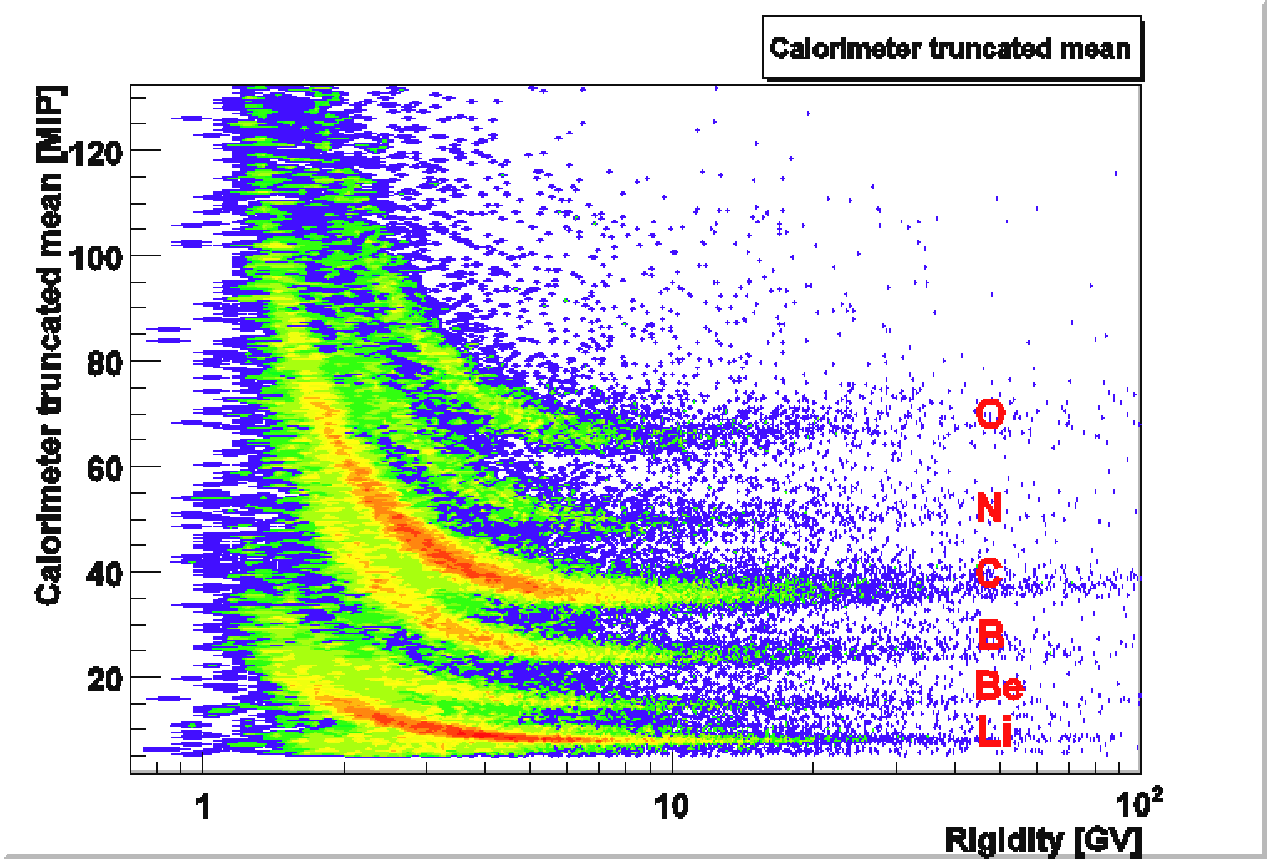
Electron energy, dE/dx, lepton-hadron separation

- 44 Si-x / W / Si-y planes (380)
- 16.3 X0 / 0.6 L
- dE/E ~5.5 % (10 - 300 GeV)
- Self trigger > 300 GeV / 600 cm^2 sr

- 36 ^{3}He counters
- ^{3}He(n,p)T; E_p = 780 keV
- 1 cm thick poly + Cd moderator
- 200 μs collection

PAMELA

Ionisation in the calorimeter ® Z identification of nuclei

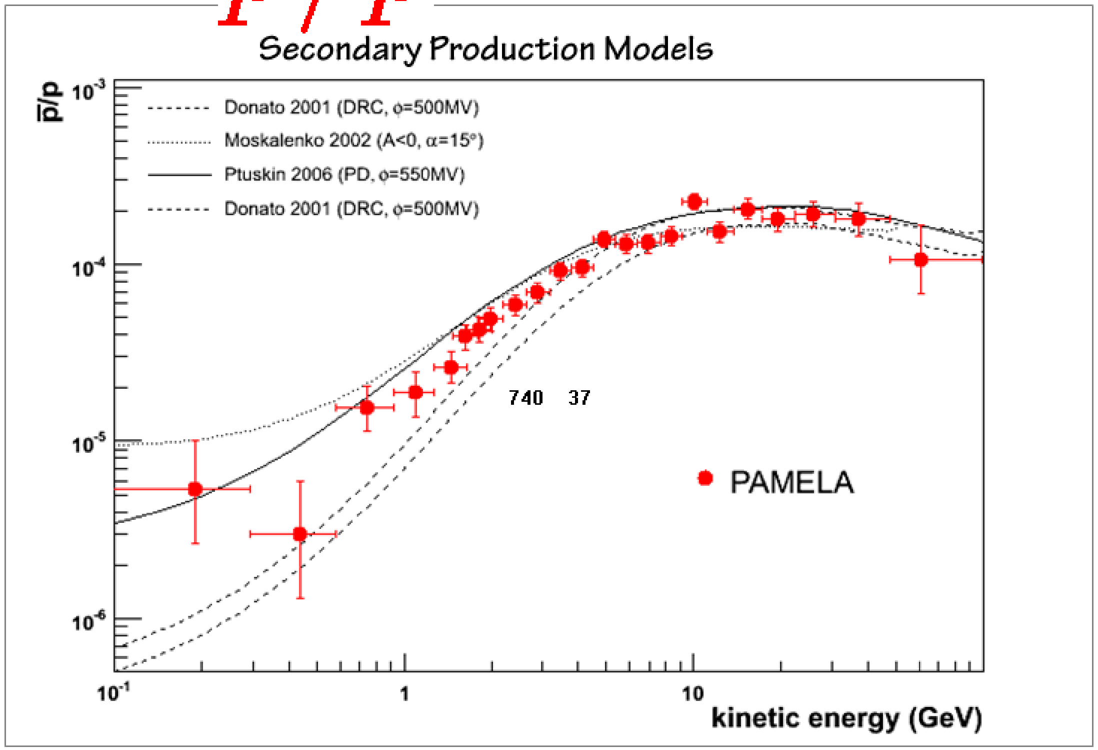


Charge identification capabilities (calorimeter)

PAMELA

Antimatter ?

$\bar{p}/p$



Antiproton to proton ratio

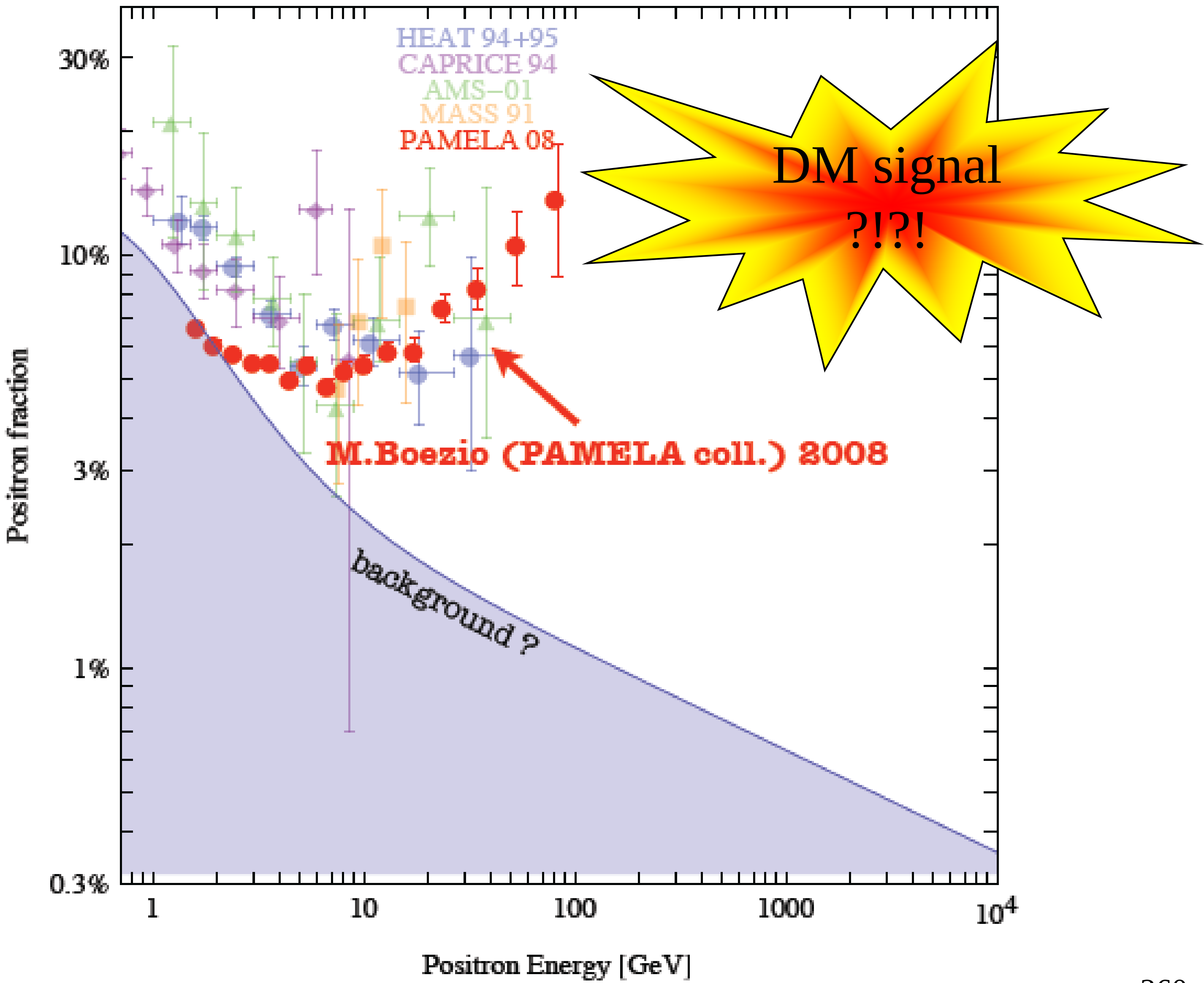
o-ph 0810.4994

PAMELA

Antimatter ?

$$\frac{e^+}{e^- + e^+}$$

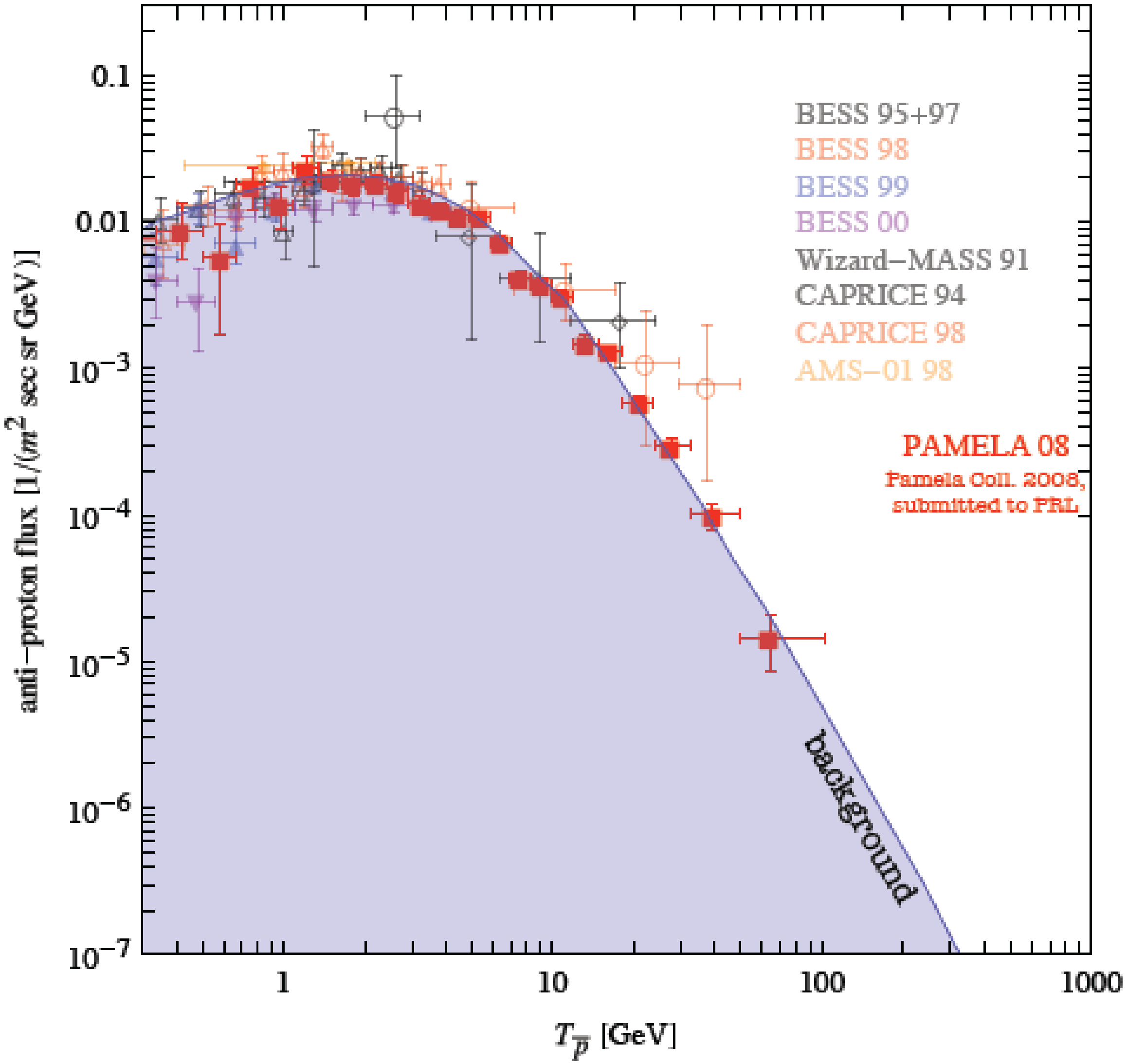
strong exces above 10 GeV
Important flux confirmed by AMS-2



PAMELA

Antimatter ?

$\bar{p}/p$
Consistant
with "background"



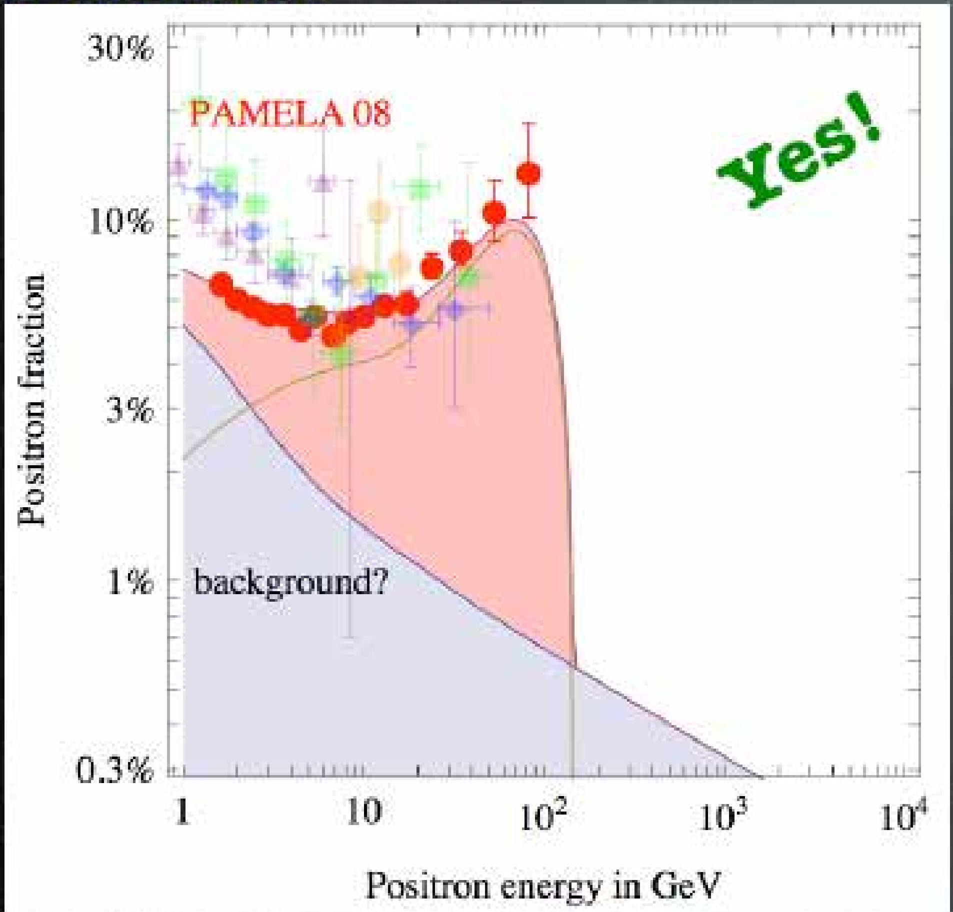
M.Cirelli, seminar @LPSC
December 2009

Results

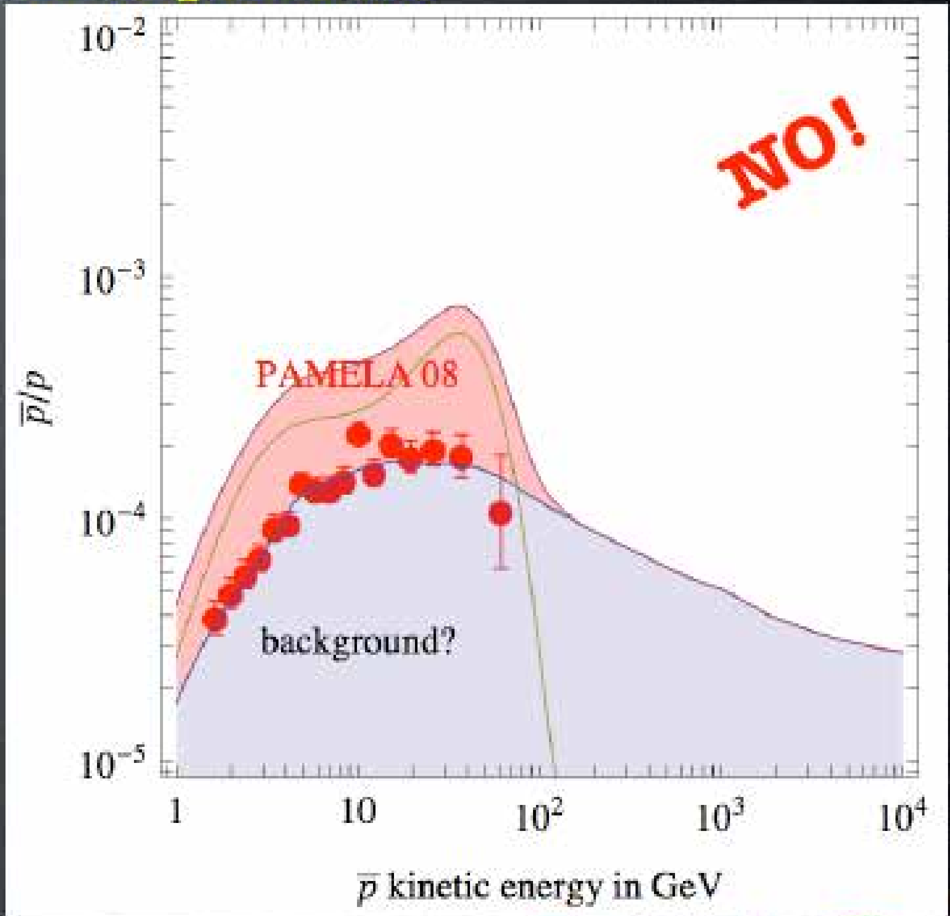
Which DM spectra can fit the data?

E.g. a DM with: -mass $M_{\rm DM} = 150\,{\rm GeV}$
-annihilation $\rm DM\ DM \to W^+W^-$
(a possible SuperSymmetric candidate: wino)

Positrons:



Anti-protons:



[insisting on Winos]

M.Cirelli, seminar @LPSC
December 2009

Results

Which DM spectra can fit the data?

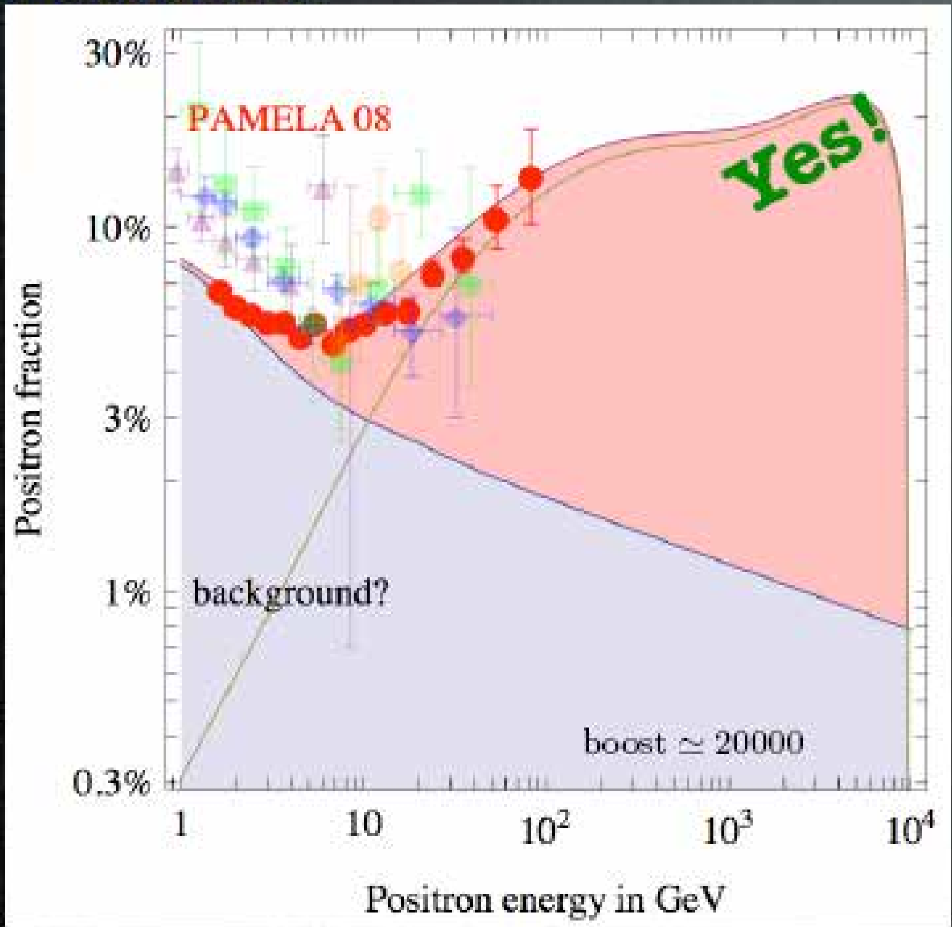
E.g. a DM with: -mass $M_{\rm DM} = 10\,{\rm TeV}$

-annihilation $\rm DM\, DM \rightarrow W^+W^-$

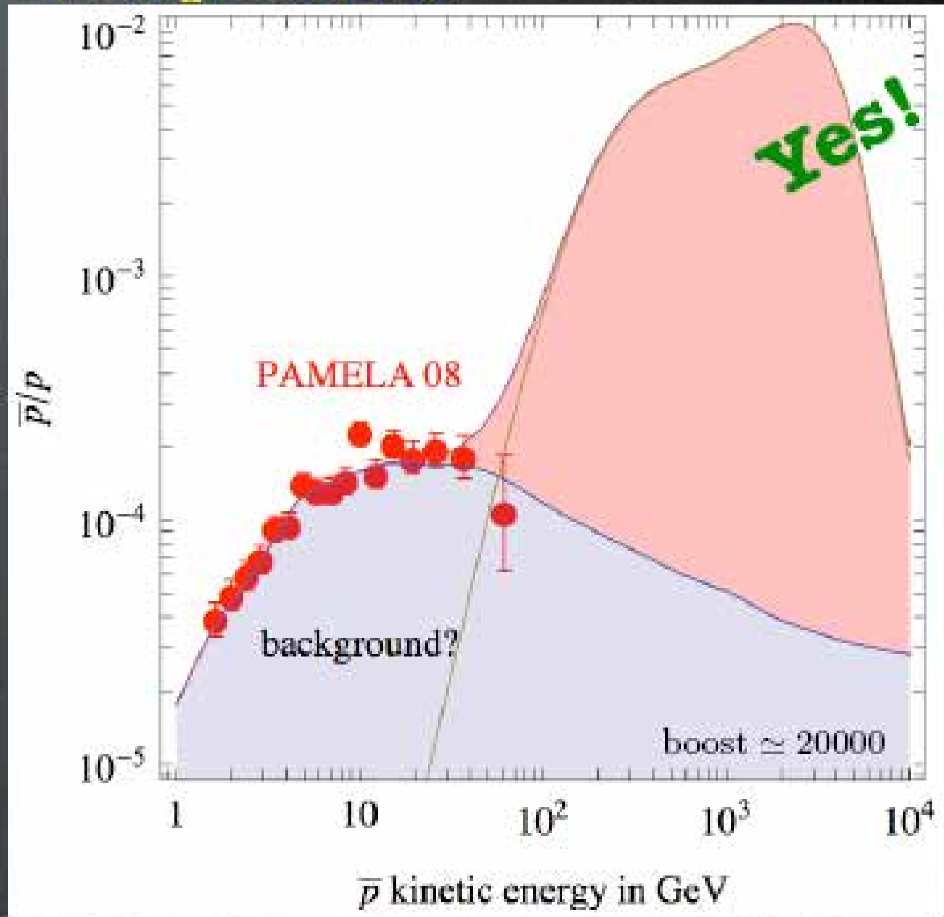
but...: -cross sec $\sigma_{\rm ann} v = 6 \cdot 10^{-22} {\rm cm}^3/{\rm sec}$

Mmm...

Positrons:



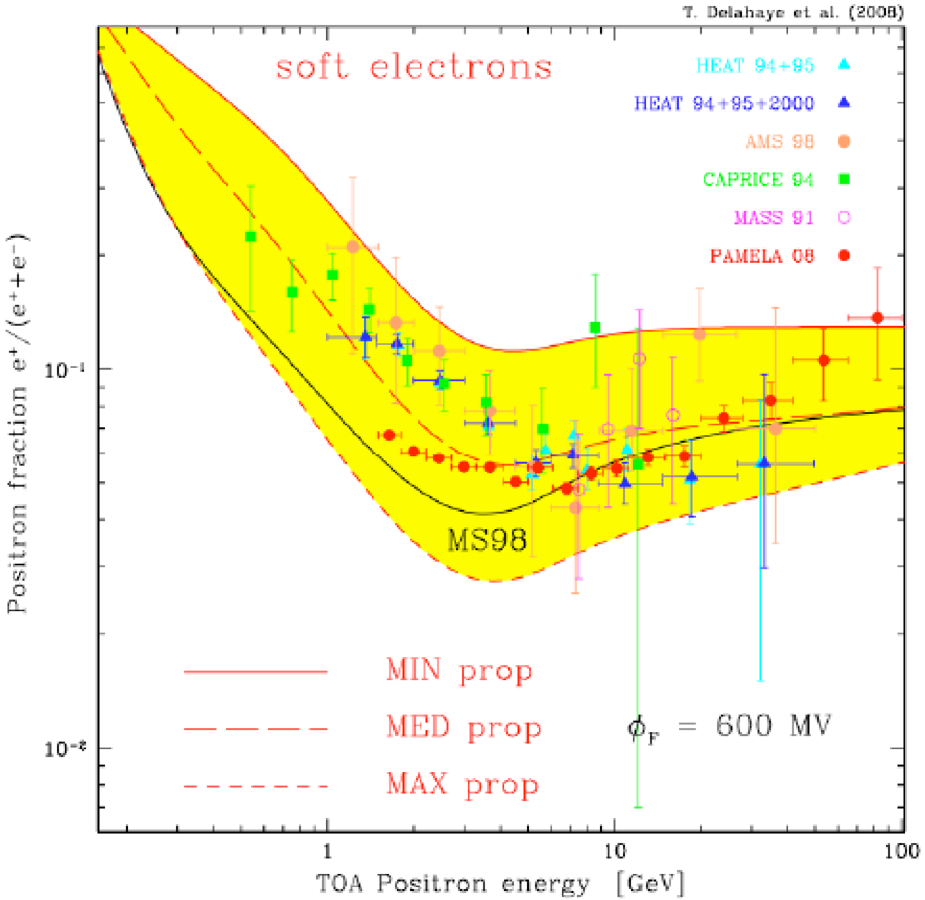
Anti-protons:



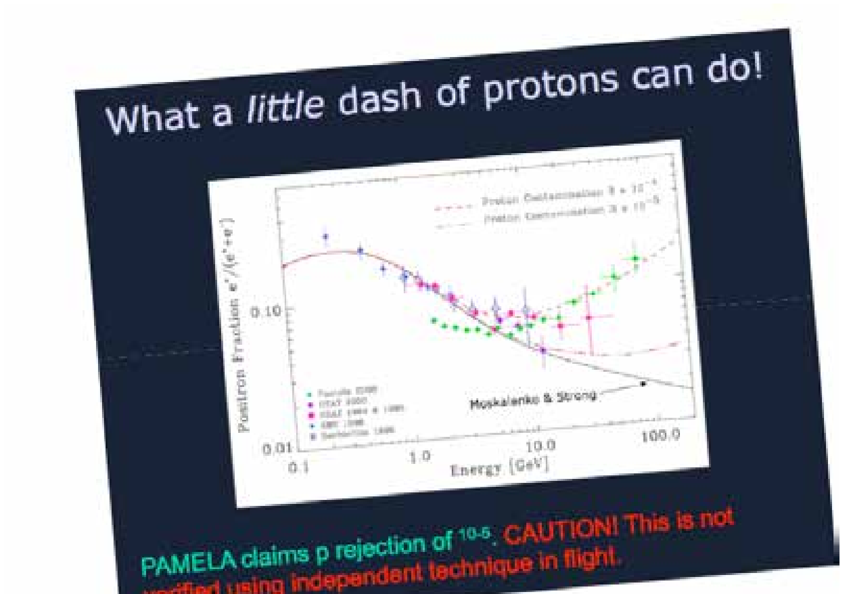
Background "measurements"

e+ background:
mainly spallation of CR on IS gas.

Baltz, Edsjo 1999
Moskalenko, Strong 1998
Delahaye et al., 0809.5268
P.Salati, Cargese 2007

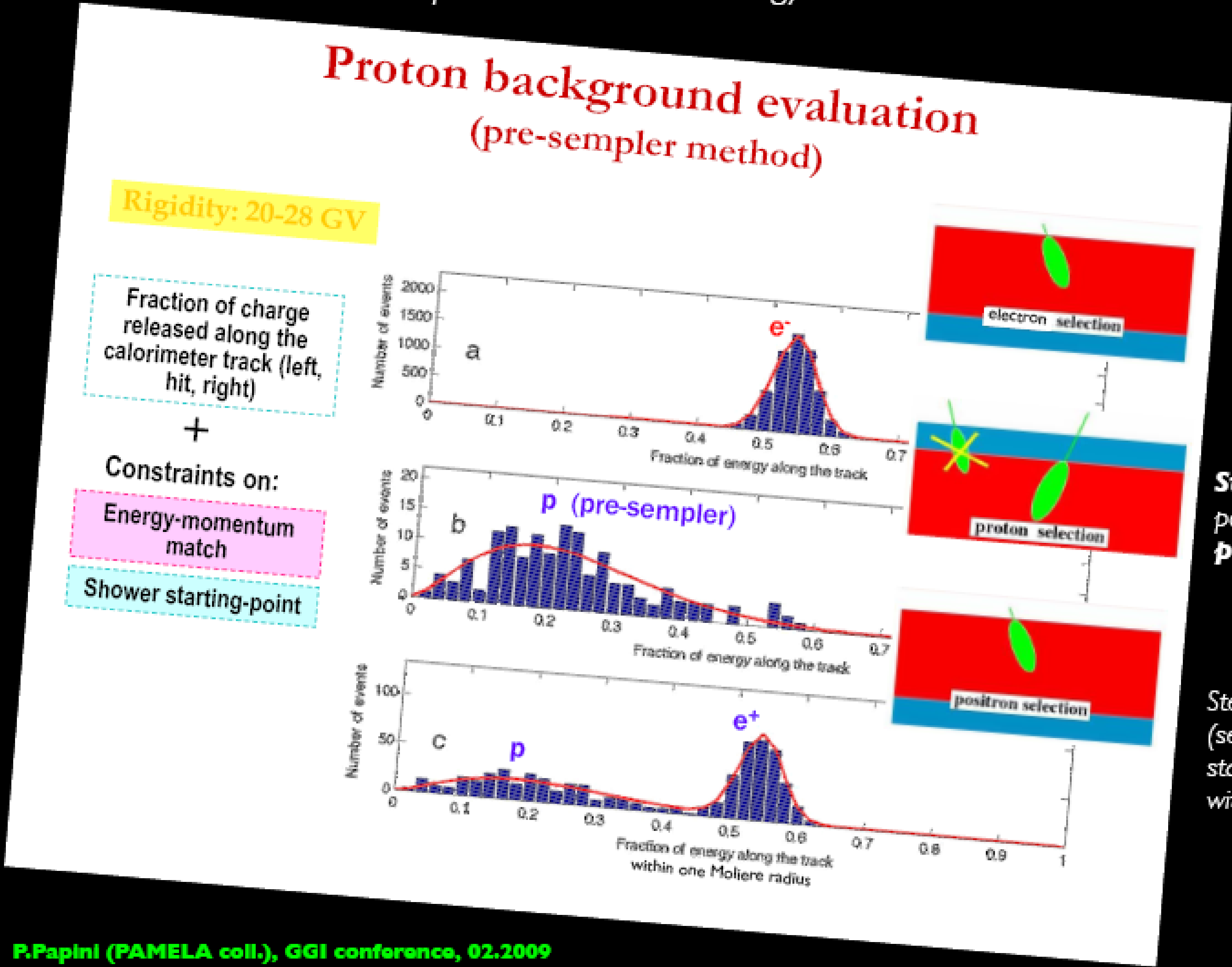


Control of instrumental effects ?



"PAMELA did do in-flight checks of the p rejection rate"

Method: in the calorimeter, leptons leave all their energy and on the top;
protons leave little energy and in the bottom.



Step 1: use the upper portion of the calorimeter to select electrons only ($\bar{p}$ negligible)

Step 2**: shower in lower portion selects **protons only

Step 3: full analysis (see that peak is statistically consistent with e^- peak of step 1)

P.Papini (PAMELA coll.), GGI conference, 02.2009

A nearby Pulsar ?

Mature pulsars: (0.01 < T < 1 Myr)

$e^{\pm}$ are confined to the pulsar wind nebula until it merges with the ISM. Merger process is fast so that pulsars can be treated as burst-like sources of $e^{\pm}$

Contribution could be from a few local pulsars
- Geminga: d = 160 pc, T = 0.37 Myr
- Monogem: d = 290 pc, T = 0.11 Myr

but pair conversion efficiency needs to be high (30 - 40%)

Contribution could be from a large number of pulsars, distant and local, with an assumed continuum distribution and injection spectrum

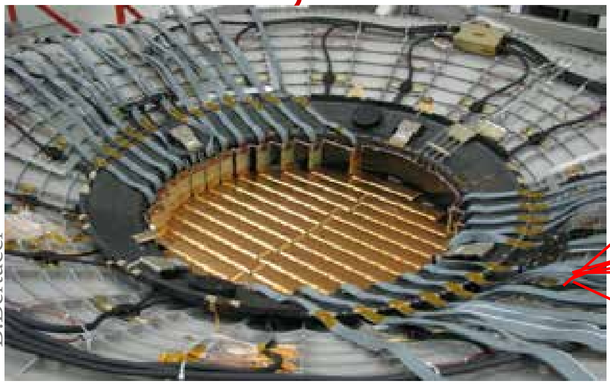
$$\frac{dN_{e^{\pm}}}{dE} \propto E^{-1.5} e^{-E/E_p}$$

Moriond 2-9 March 2013

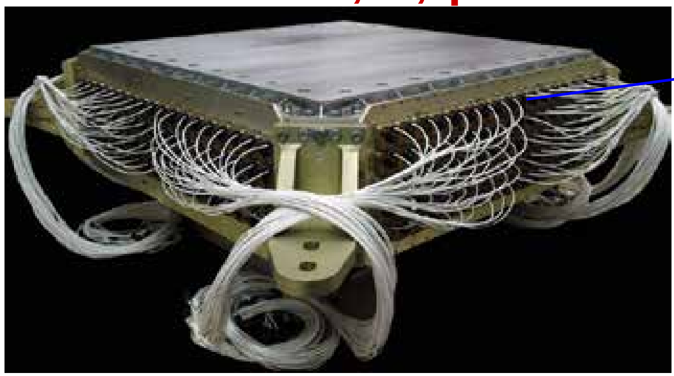
A precision, multipurpose spectrometer up to TeV

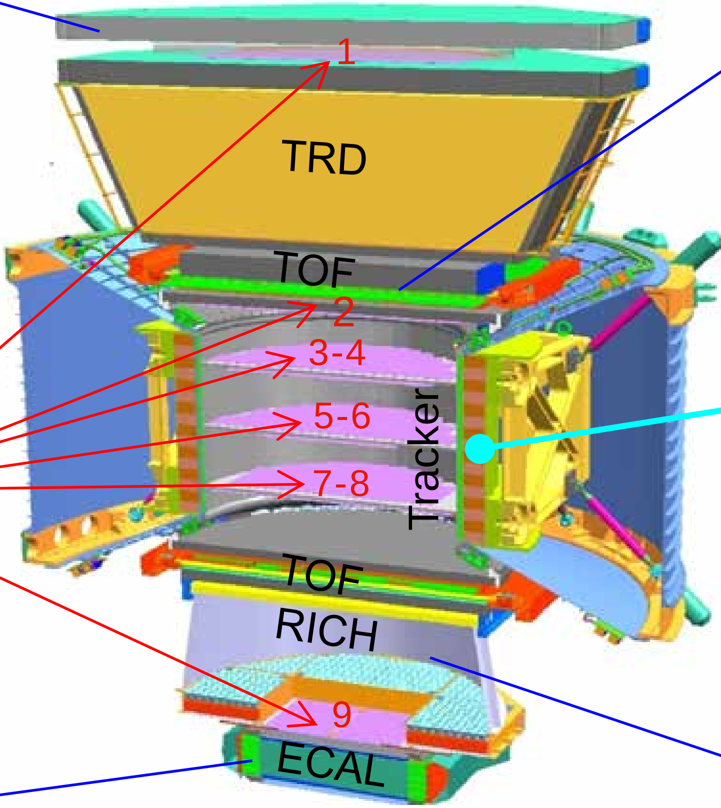
TRD
Identify e^+, e^-

Silicon Tracker
Z, P



ECAL
E of e^+, e^-, γ

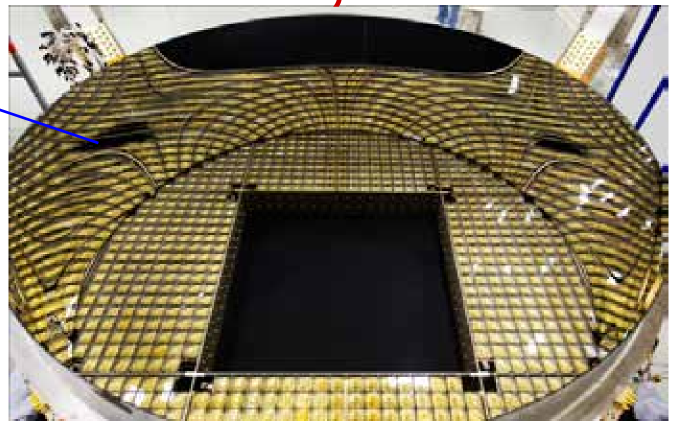




TOF
Z, E

Magnet
±Z

RICH
Z, E

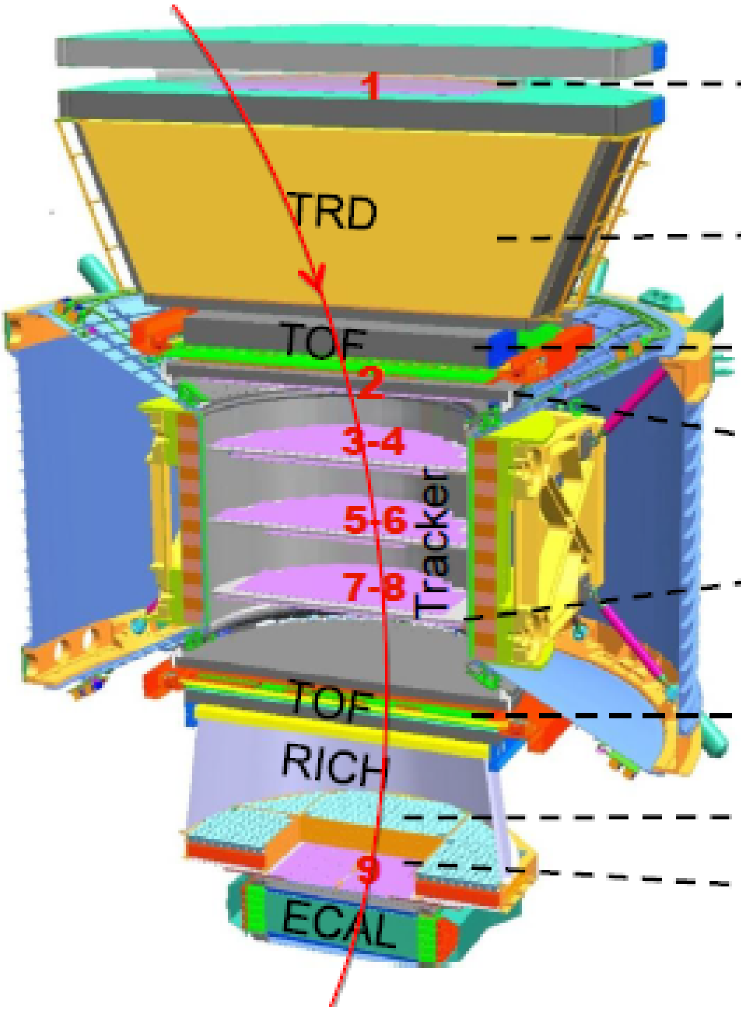


***Z, P* are measured independently by the Tracker, RICH, TOF and ECAL**

B.Bertucci

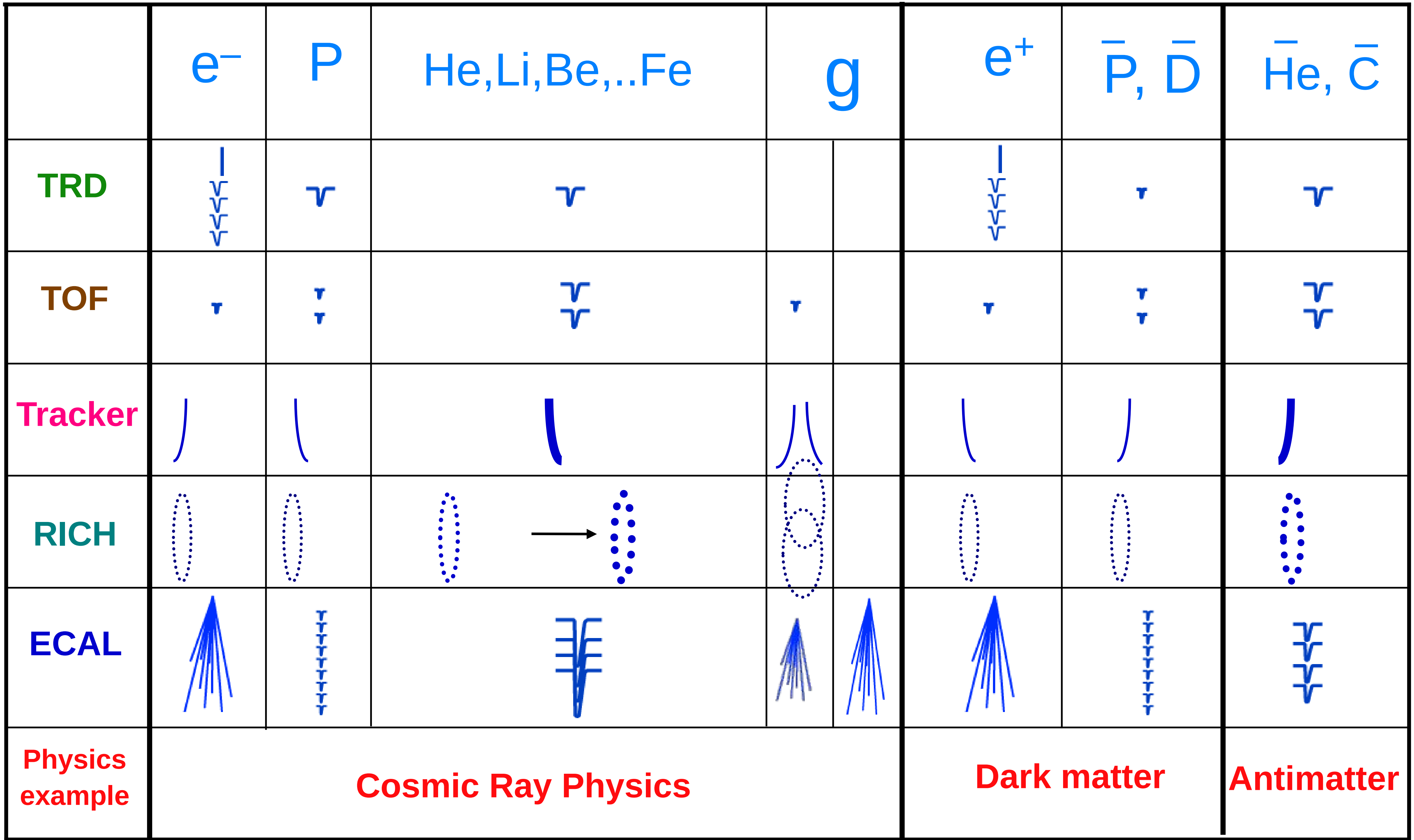
AMS charge identification

AMS: Multiple Independent Measurements of the Charge (|Z|)

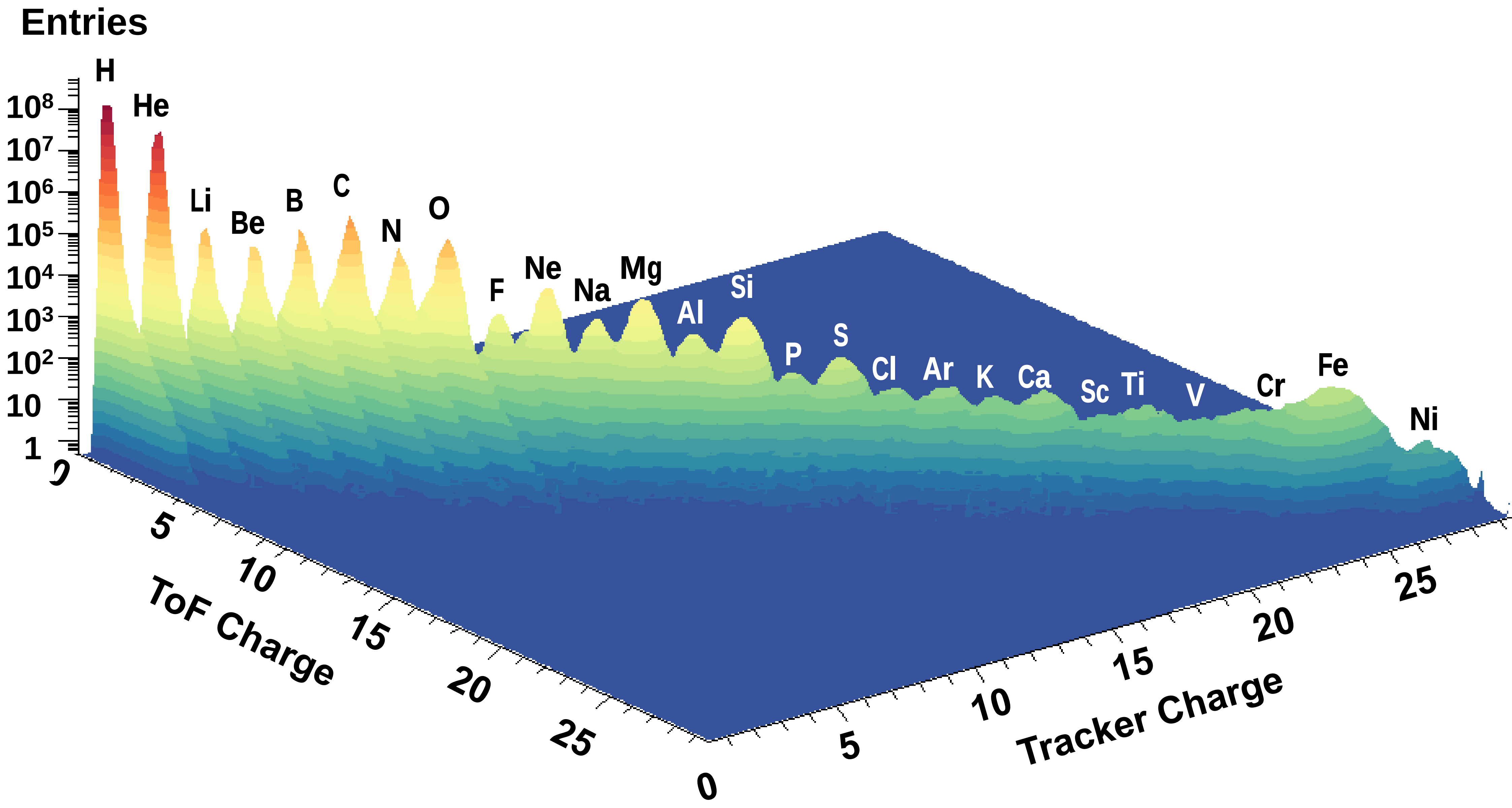


	Carbon (Z=6) ΔZ (cu)
1. Tracker Plane 1	0.30
2. TRD	0.33
3. Upper TOF (1 counter)	0.16
4. Tracker Planes 2-8	0.12
5. Lower TOF (1 counter)	0.16
6. RICH	0.32
7. Tracker Plane 9	0.30

Full coverage of anti-matter & CR physics



AMS Nuclei Measurement on ISS



Charged particles (1 TeV → few 100 TeV)

Balloon Experiments

- Very high altitude flights: 40 km ´ 3,9 g cm^{-2}
- Load: up to ≈ 270 kg and 10 m^3
- Many former flights:

JACEE (USA + Japon) ; RUNJOB (Russie + Japon) etc.

using emulsion chambers !

- Recently: Ultra Long Duration Balloon (ULDB) Flights
60-100 jours (NASA, Antarctic) → the CREAM experiment

(Cosmic Ray Energetics and Mass)

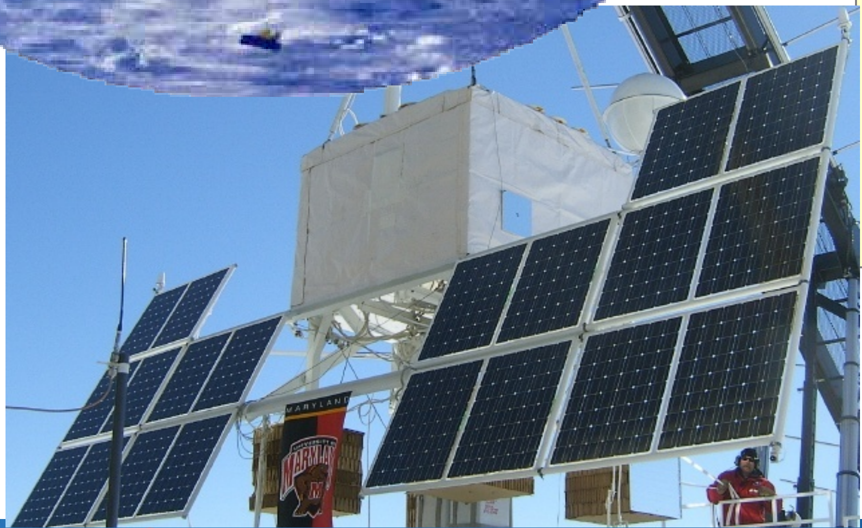
CREAM

Ultra Long Duration Balloon

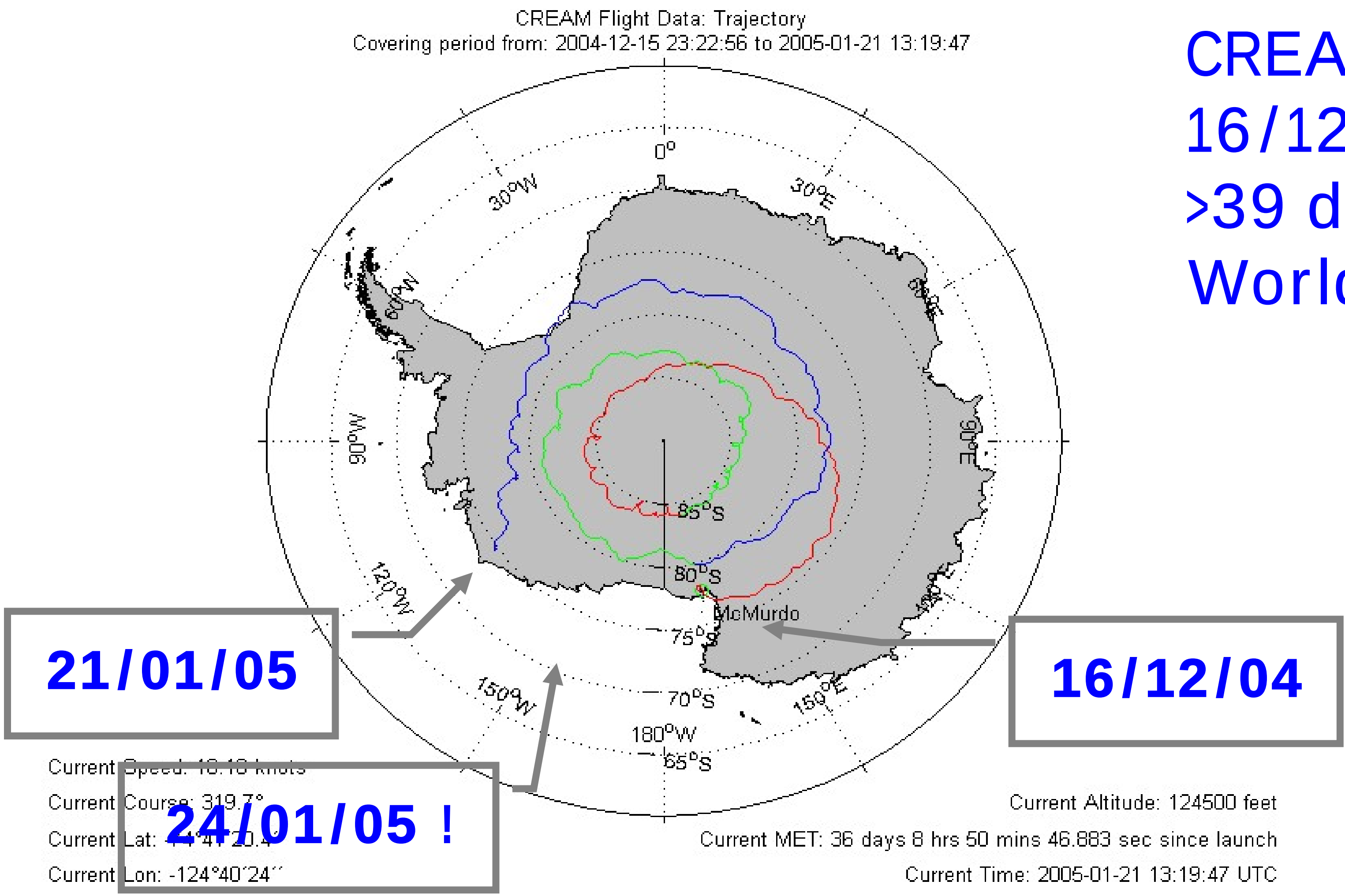
ULDB Proj., Adv.Sp.Res33,1633(2004) :

NASA project to develop

- Flight of < 100 days
- Payload · 2 tons
- Alt 33000 meter
- CREAM n° 1 : 2006 (2005/LDB)

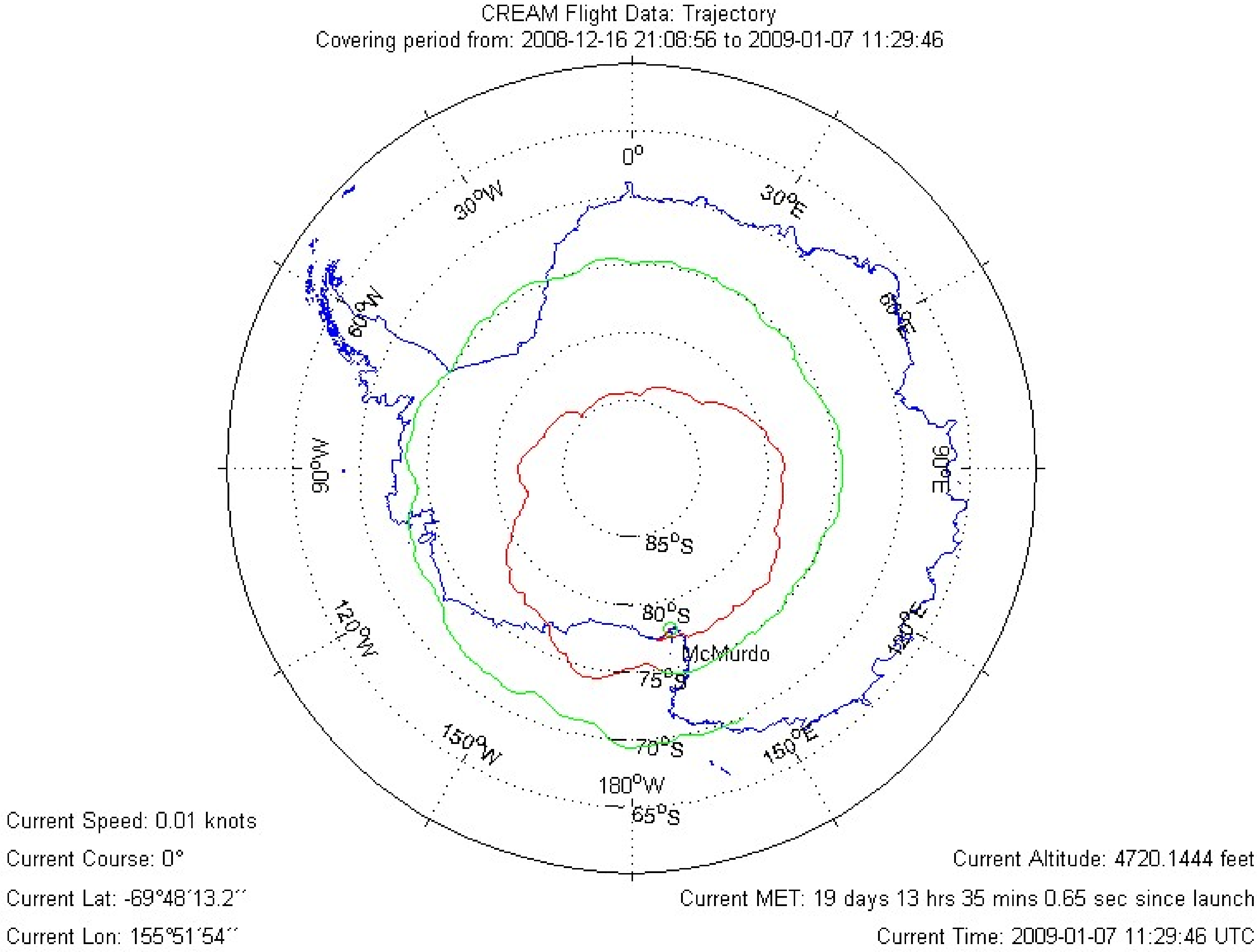


CREAM 2004



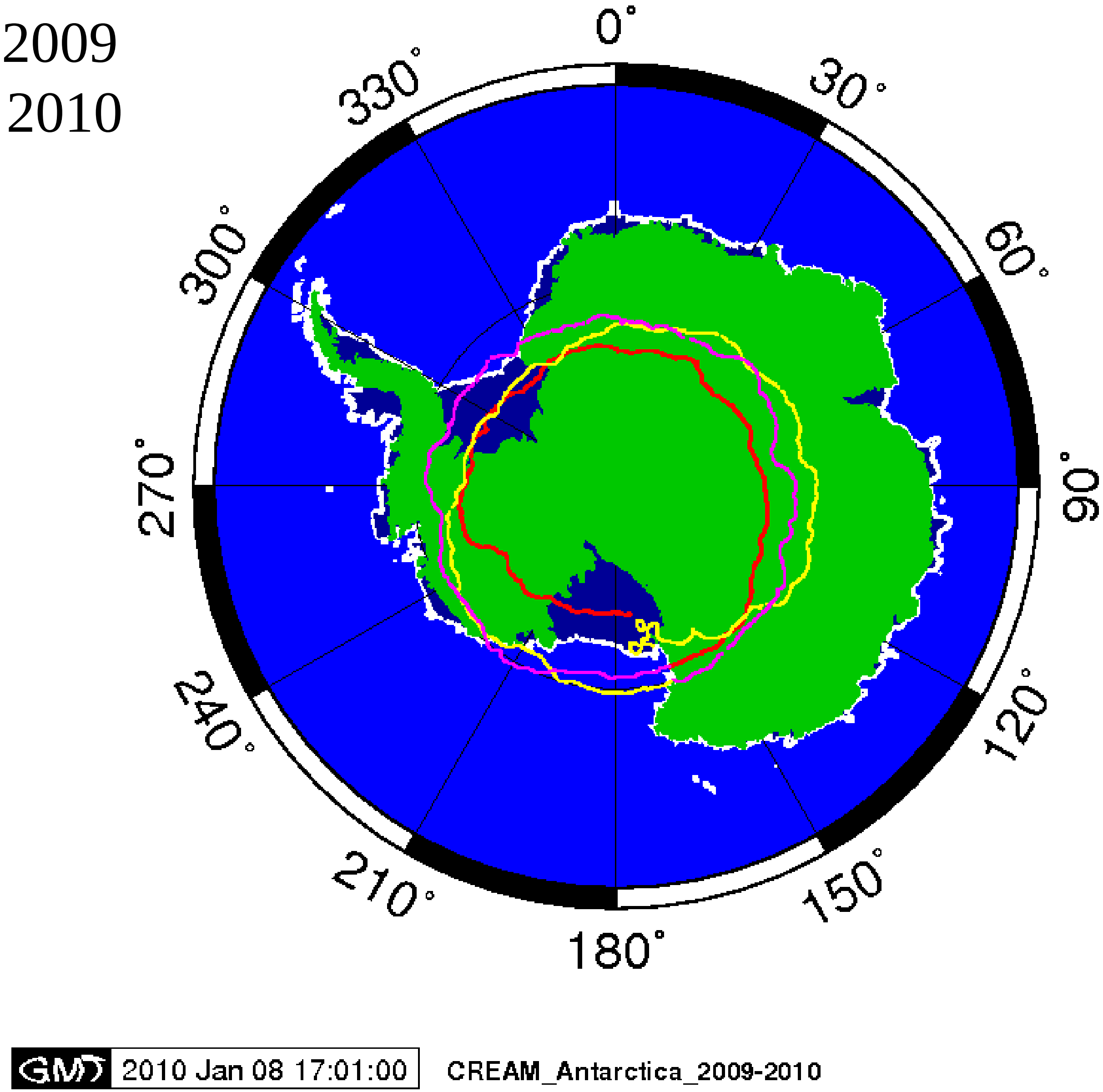
CREAM I:
16/12/04 - **/01/05
>39 days flight
World Record

CREAM 2008


CREAM Flight Data: Trajectory
Covering period from: 2008-12-16 21:08:56 to 2009-01-07 11:29:46
0°
30°E
60°E
90°E
120°E
150°E
180°W
150°W
120°W
90°W
60°W
30°W
85°S
80°S
75°S
70°S
65°S
McMurdo
Current Speed: 0.01 knots
Current Course: 0°
Current Lat: -69°48´13.2´´
Current Lon: 155°51´54´´
Current Altitude: 4720.1444 feet
Current MET: 19 days 13 hrs 35 mins 0.65 sec since launch
Current Time: 2009-01-07 11:29:46 UTC

CREAM 2009

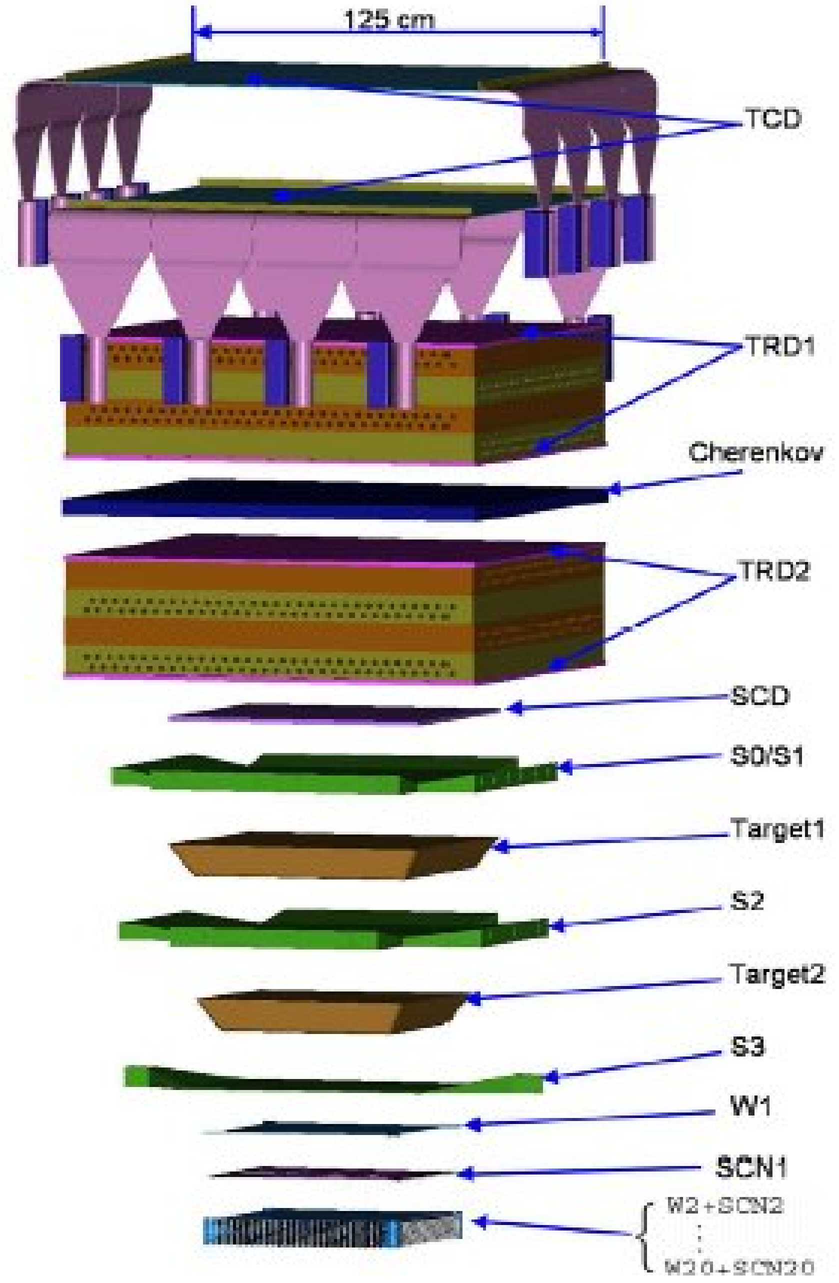
1 décembre 2009
au 8 janvier 2010



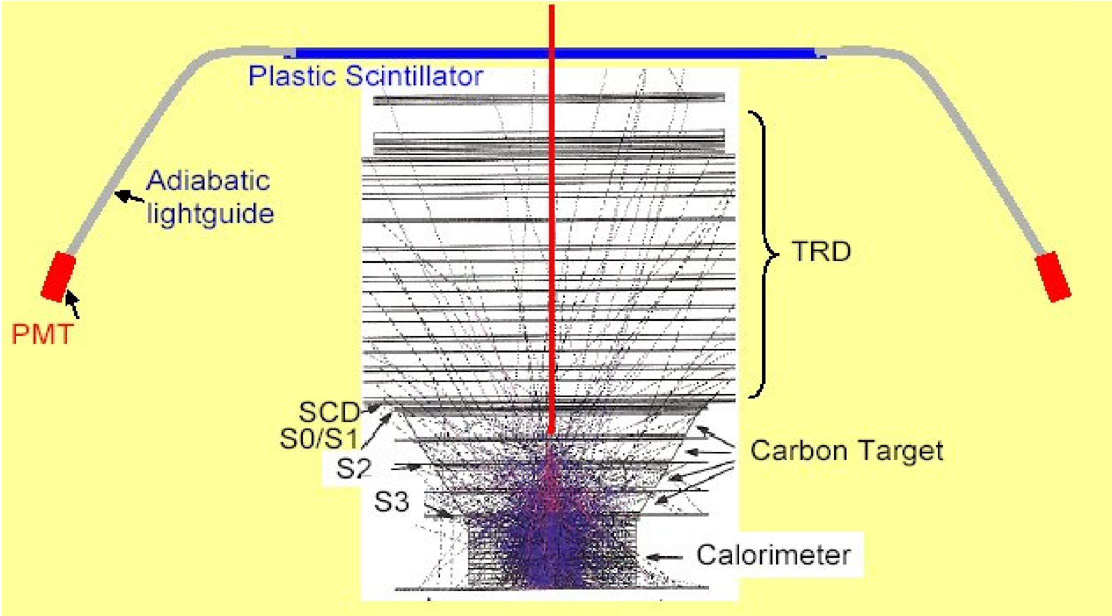
CREAM

Cosmic Ray Energetics and Mass

- Objectives :
  CR composition and spectrum of the different elements
  (from TeV to ~500 TeV)
- Acceptance : 2,2 m^2 sr
- Energy measurement:
  - Calorimeter 20 X_0 (W + scint. fibres)
  - Transition Radiation Detector
- Identification :
  - TRD
  - Cherenkov detector "CHERCAM" similar to AMS-2



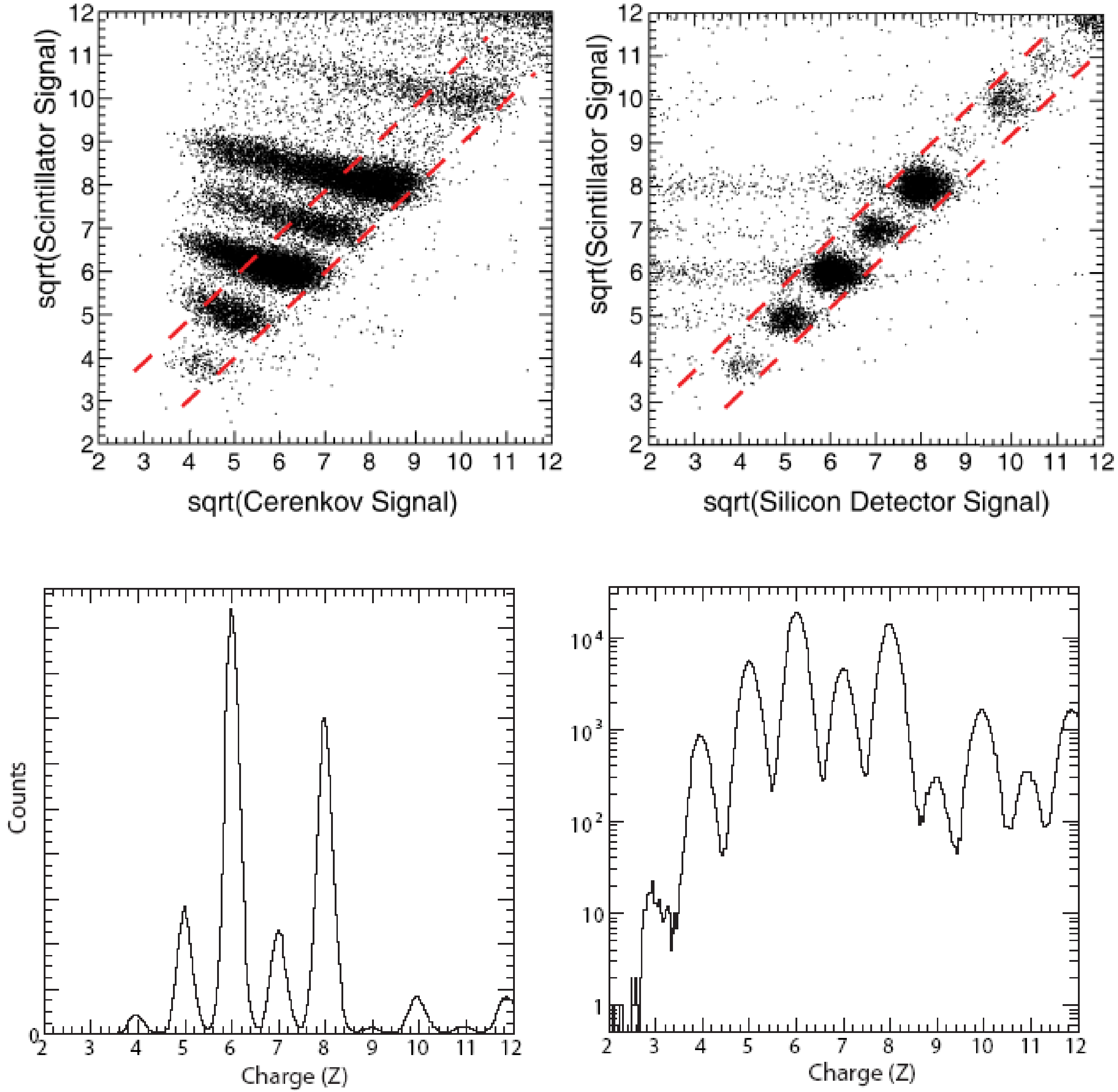
CREAM experiment



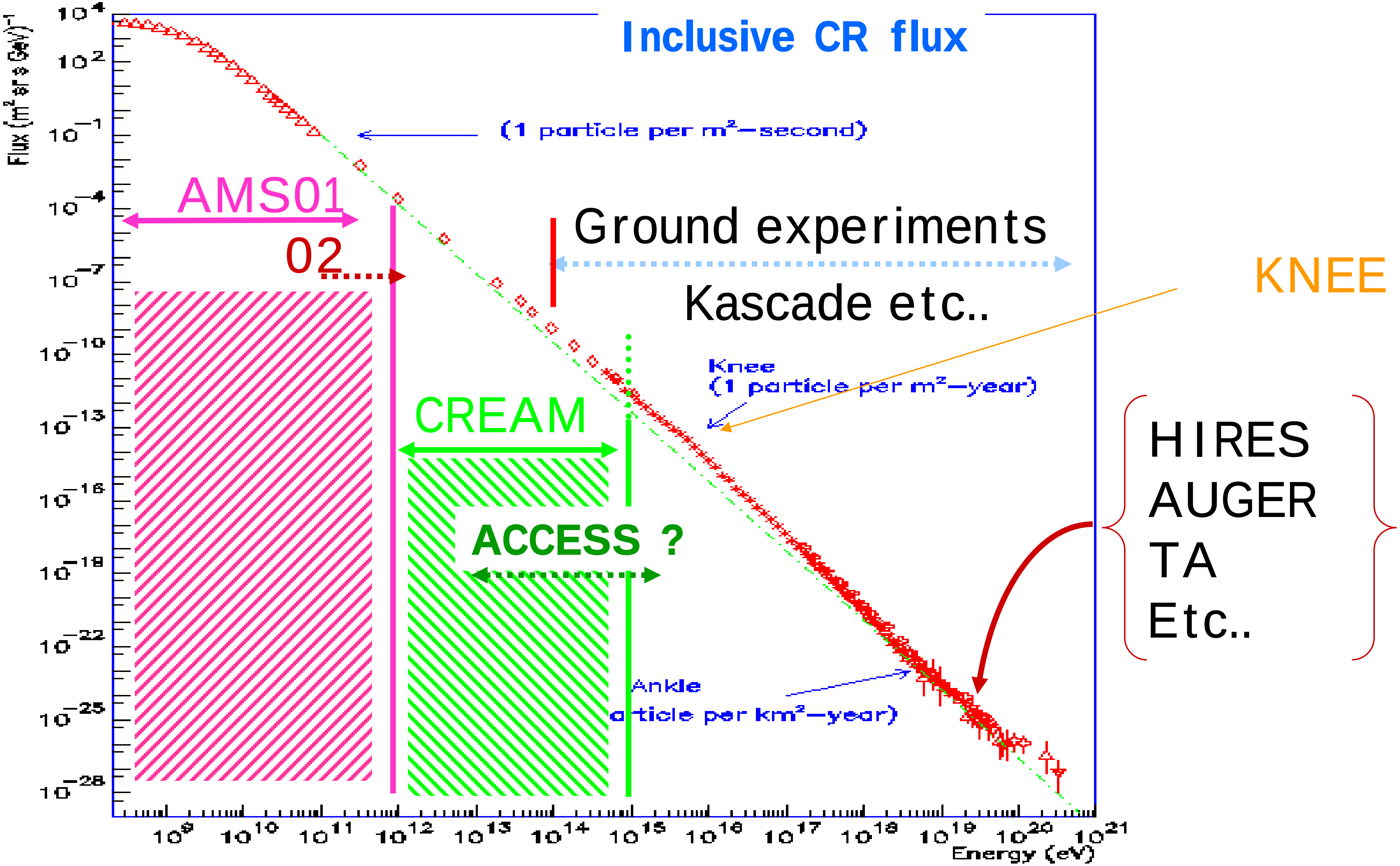
- At TeV energies, the interaction of CR in the calorimeter induces many backscattered secondary particles that one have to veto.

- The "CHERCAM" cherenkov solves this problem by measuring accurately the time of any through going particle as well as achieving a precise charge measurement (± 0,3 e)

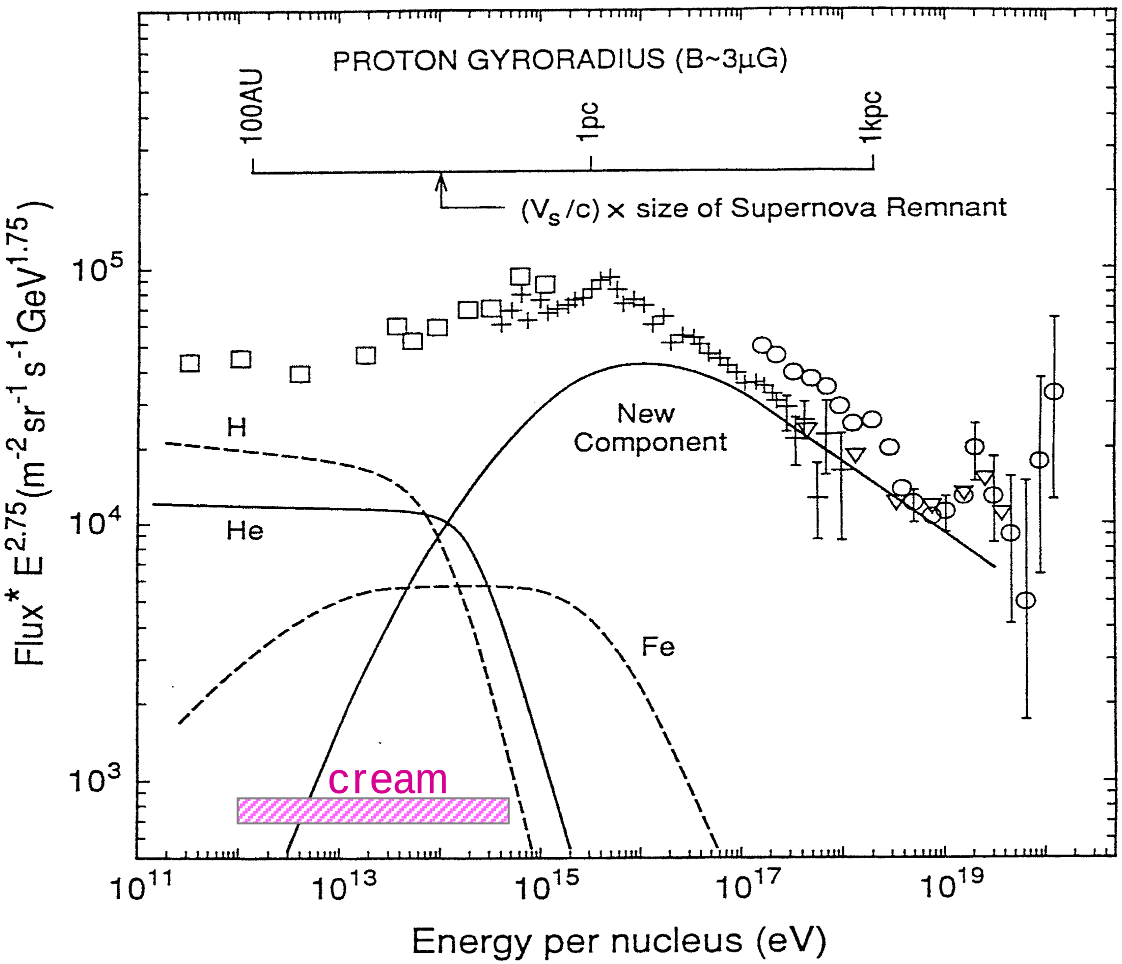
CREAM experiment



Contexte expérimental



The knee



- Is the knee due to:

- Acceleration mechanisms

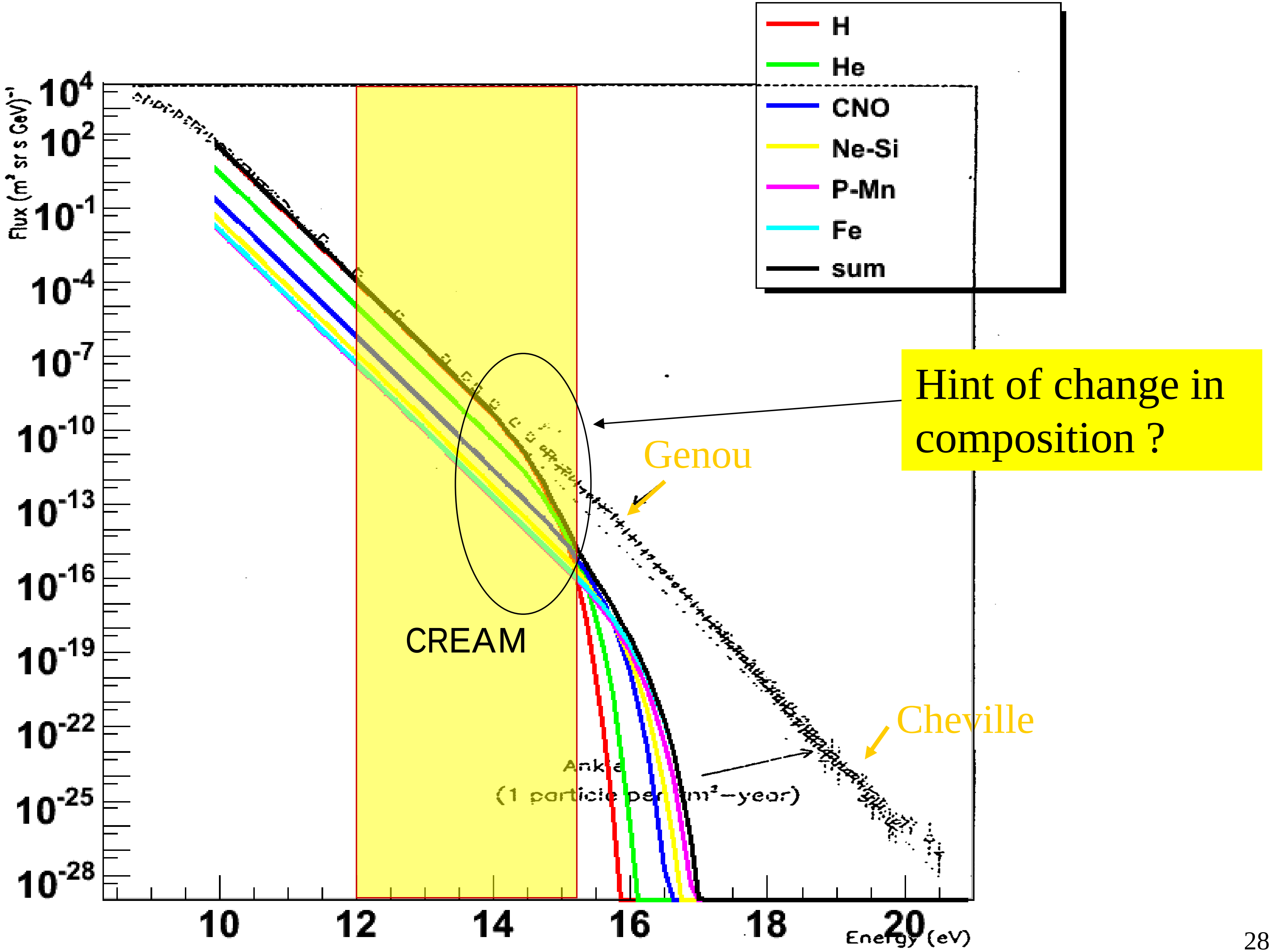
or to changes :

- in propagation?
- in CR sources?
- in interaction properties (threshold) ?

→ A diffuse SNR shock acceleration with E_{max} implies a change in composition around ~10^{14} eV.

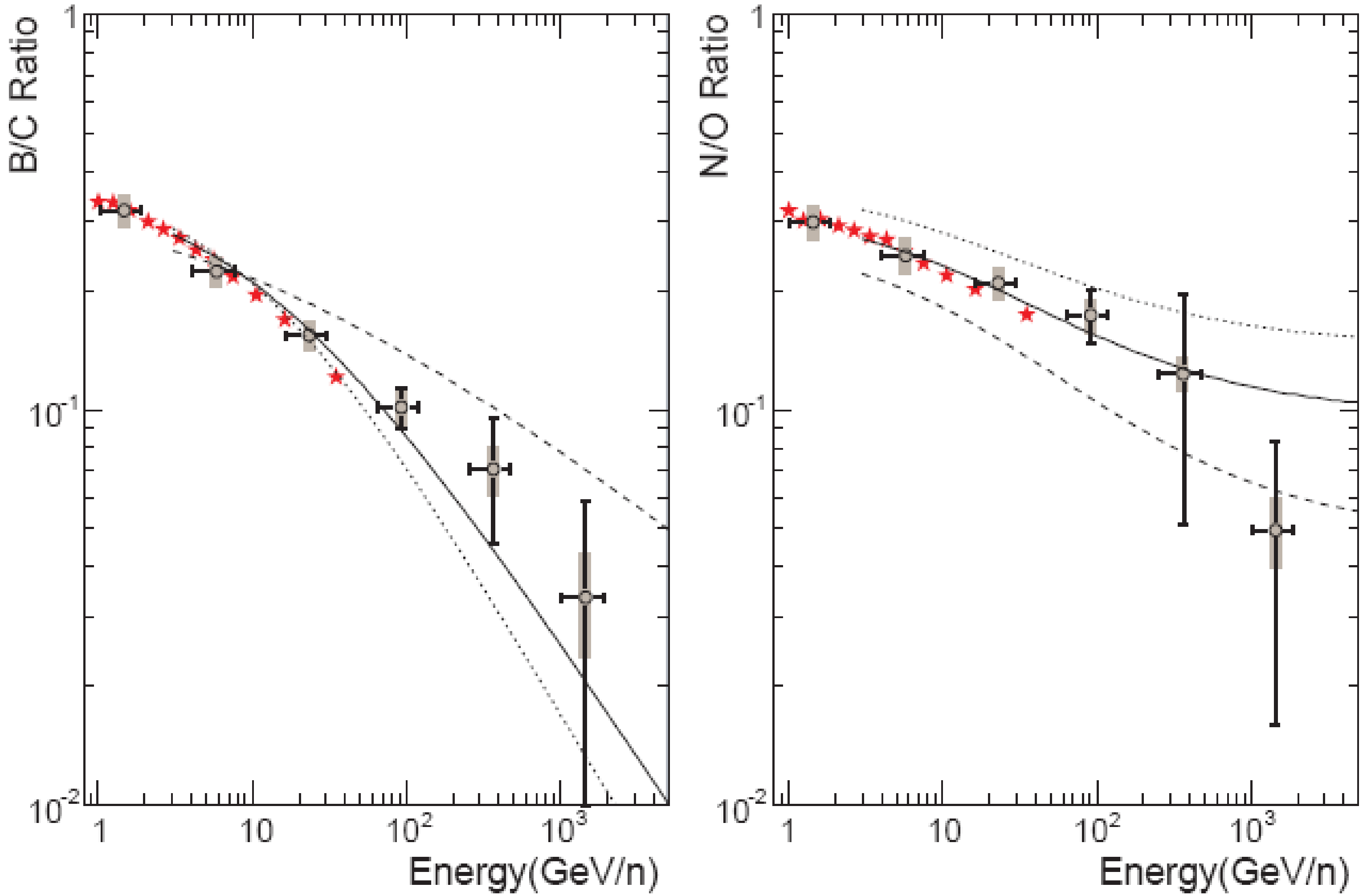
SNR energy limit: $E_{max} \sim Z \bullet 10^{14}$ eV

CR SM with Emax / Z



Galactic CR propagation with CREAM

Secondary/Primary ratio



Compatible with propagation models with escape terme in $E^{0.6}$

AIR SHOWERS DEVELOPMENT MODELS

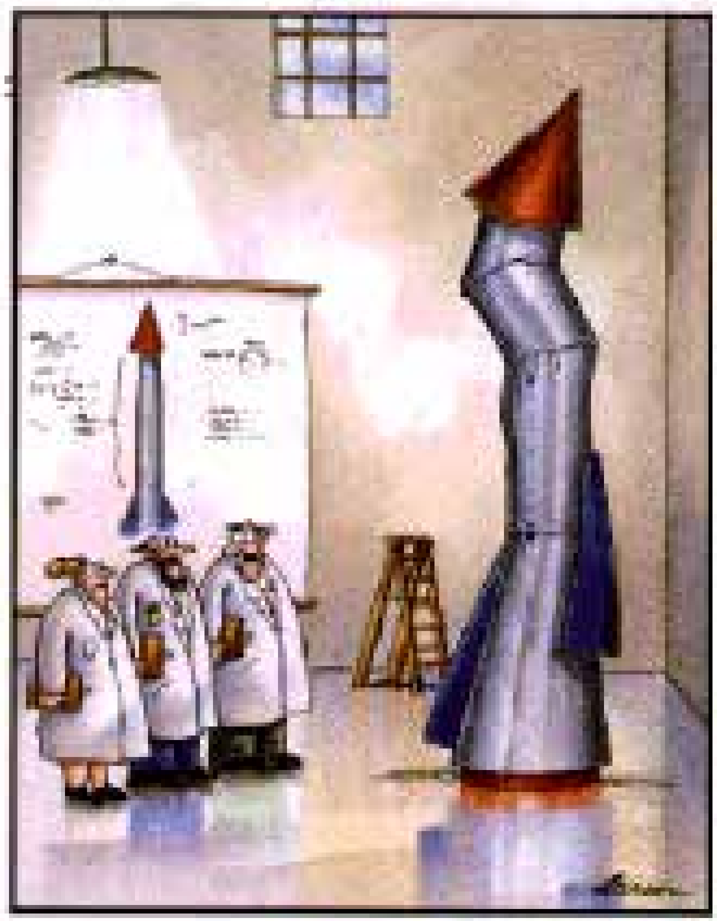
A peek above the knee !

To measure the inclusive spectrum at the knee, one nees a $10m^2$ exposed during 10 years !

The realistic experimental limits are:

- For satellites ~ $1m^2$ (sr) during ~few years
- For balloons, ~ $10m^2$ (sr) during n´ 30 jours

® $E < 10^{15}$ eV



It's time we face reality my friends, we should keep to ground detectors !

Extensive Air Showers: the phenomenon and the observables

- The large shower of secondary particles induced by the interaction of a primary CR in the upper atmosphere can be detected on an extensive area
  → large effective surfaces to fight against low flux at E ≥ 1000 TeV

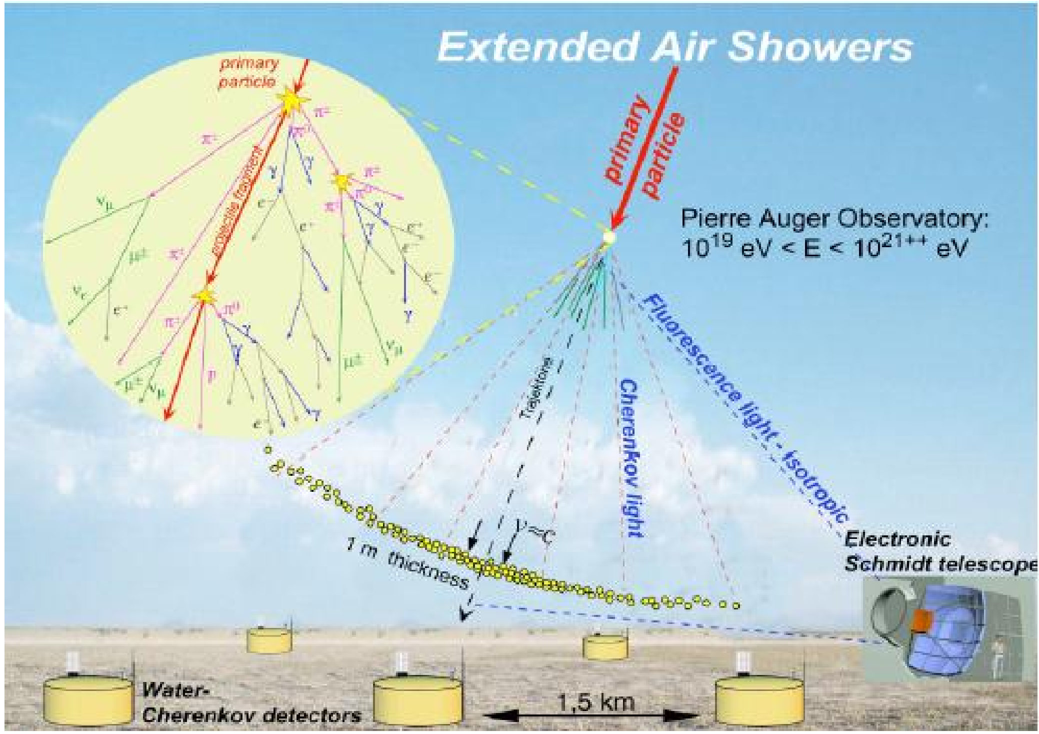
- Atmosphere used as an calorimeter
  (~1000 g cm^{-2} at sea level for a vertical shower)

- From the observables, one aims at measuring:
  - Incident direction;
  - Primary energy E_0 ;

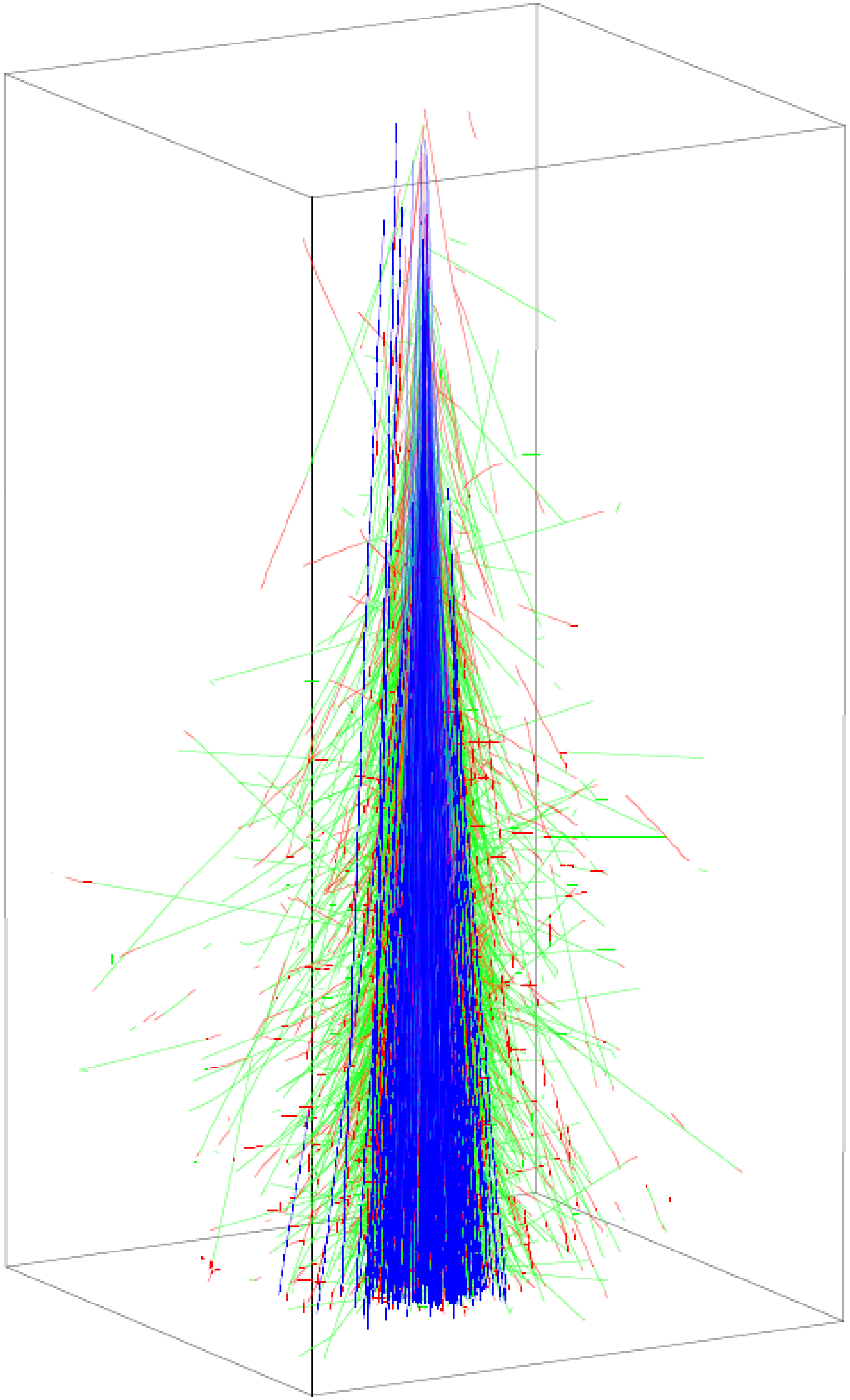
  and if possible, get access to the nature of the primary particle :
  - distinction γ-hadron ;
  - distinction light nuclei (p, He) - heavy nuclei(Fe)

p or nucleus + N or O nucleus → hadronic cascade

- Hadronic component: nuclear fragments, nucleons, mesons π, K, etc.
- Electromagnetic component: induced by $\pi^0 \to \gamma\gamma$ and other radiative decays
- Muonic component: induced by decays of $\pi^\pm$ and $K^\pm$
- Atmospheric Neutrinos issued from s $\pi^\pm$ $K^\pm$ and $\mu^\pm$ decays



Primary electrons and ° induce an electromagnetic shower consisting mainly of secondary electrons, positrons and ° (muon poor)



Shower development

« des giboulées d'électrons »

Rayon cosmiques
par Pierre Auger
1941 PUF

A 10^{19} eV shower

10^{11} particles
at sea level

Photons + electrons (99%),
muons (1%)

Ground observables

- Secondary particles reaching ground

  As a function of the primary energy and of the altitude:

  - Residual Hadrons (nuclear fragments): not numerous (>11 ¸$_{int}$).
  - $e^{\pm}$: the more numerous at shower development maximum.
  - $\mu^{\pm}$: most reach ground and may penetrate deep underground.
  - ° secondaries : may be detected at ground level via $e^{+}e^{-}$ pair conversion (e.g. Cherenkov effect in water).

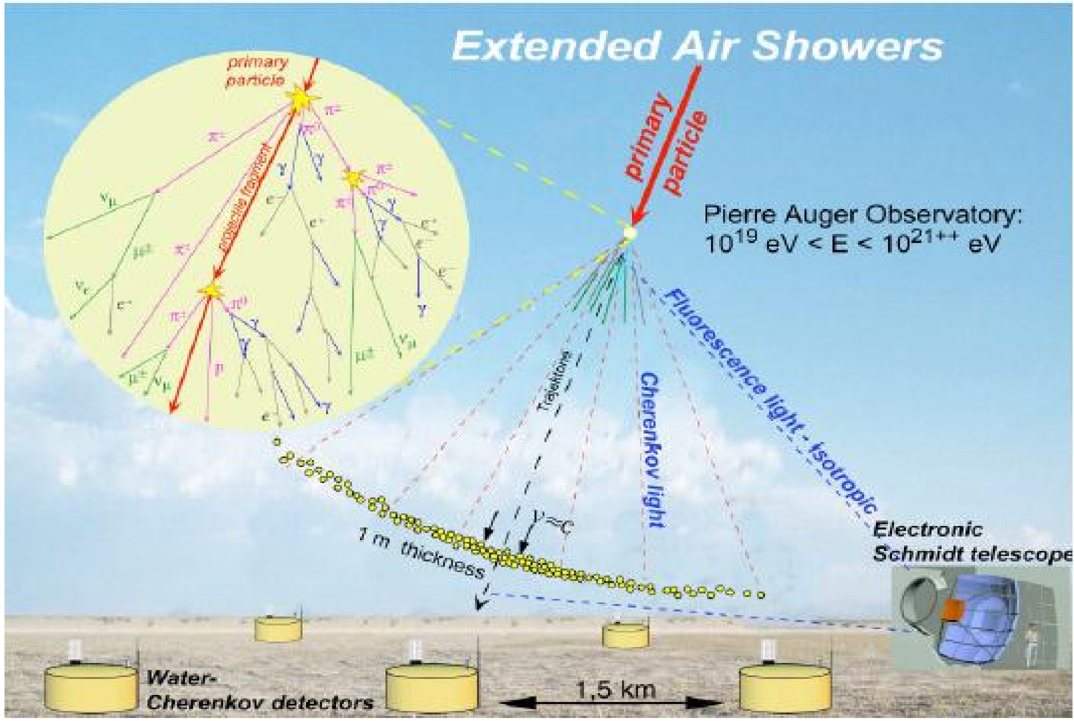
- Photons (visible, UV) emitted along the trajectories of charged particles (Cherenkov effect, N_2 fluorescence) during the shower development

  → Calorimetric 3D information !

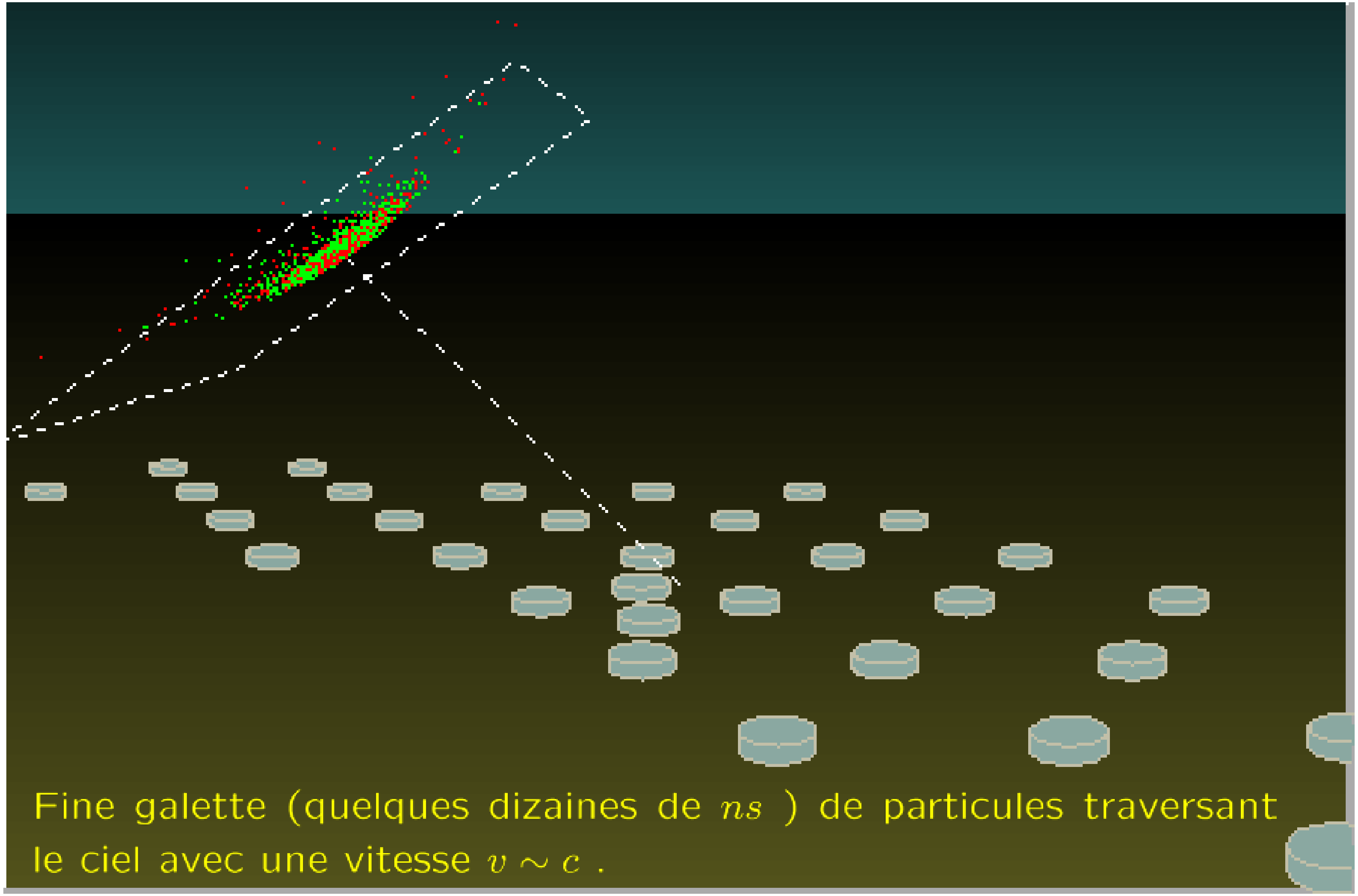
- Radio emission by the shower particles in the geomagnetic field or by the induced plasma.

Temporal aspects

- A light speed moving "pancake" of charged particles.
- This front is more or less curved depending on the shower development stage.
- The front thickness (~ 10 m) induce as signal time spread in each detector.
- The arrival time differences at ground on the sampling detectors → arrival direction (Δθ ≈ 1°).
- The Cherenkov light front (forward emission) is thinner (~m) than the charged particle front → well defined timing.



Time structure



EM shower
Longitudinal development

- Mean number of particles (e^+,e^- or °) crossing a plan ? to the shower axis after a slant depth t (in units of X_0).

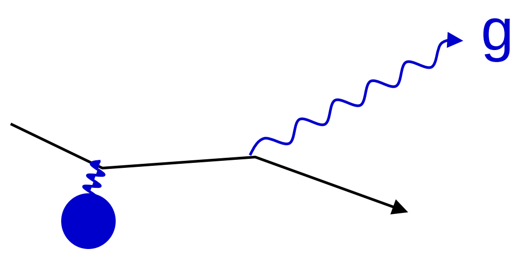
- As long as the ionization losses are small wrt radiation losses (bremsstrahlung and pair prod)
the number of particle increase exponentially.

- When the mean energy per particle decreases below the critical energy (E_c ¼ 84,2 MeV in air), the number of particle decreases
(shower extinction phase).

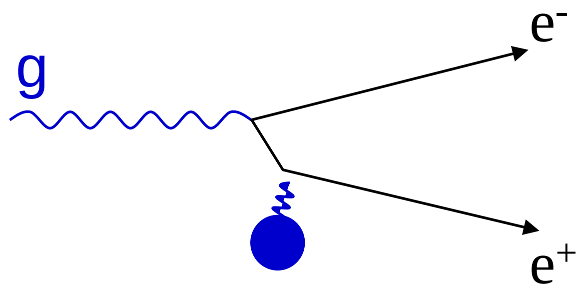
- At the transition between the two phases, (maximal development), the mean energy is equal to the critical energy.

Radiative processes ($E > E_c$)

Bremsstrahlung :



Pairs production :



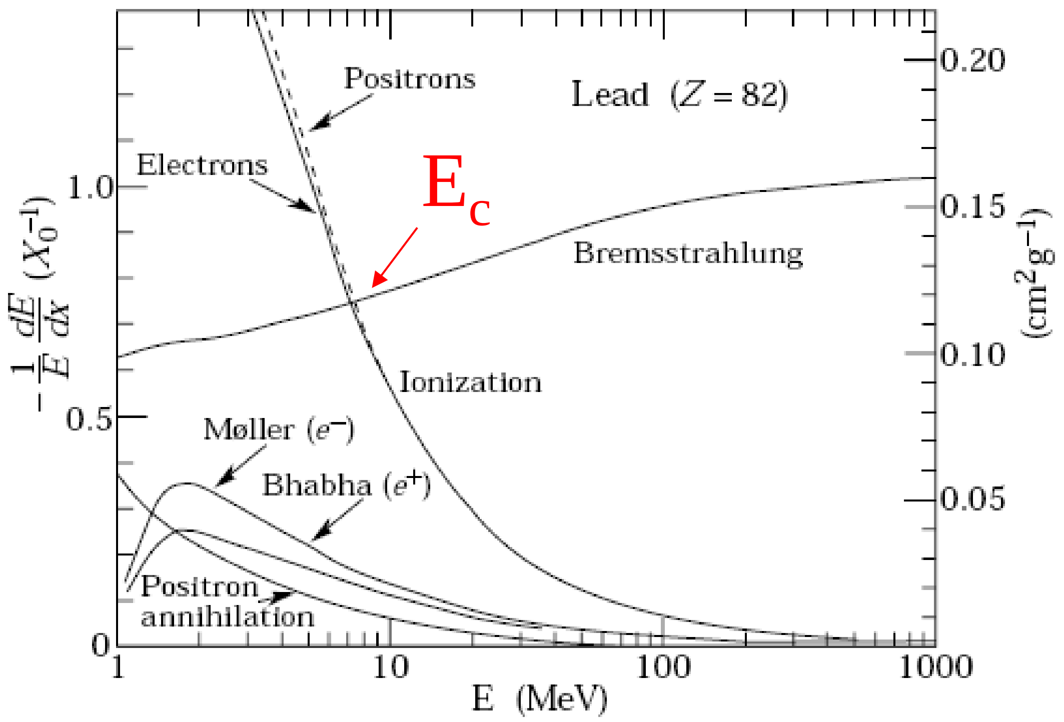
Radiation length X_0 :

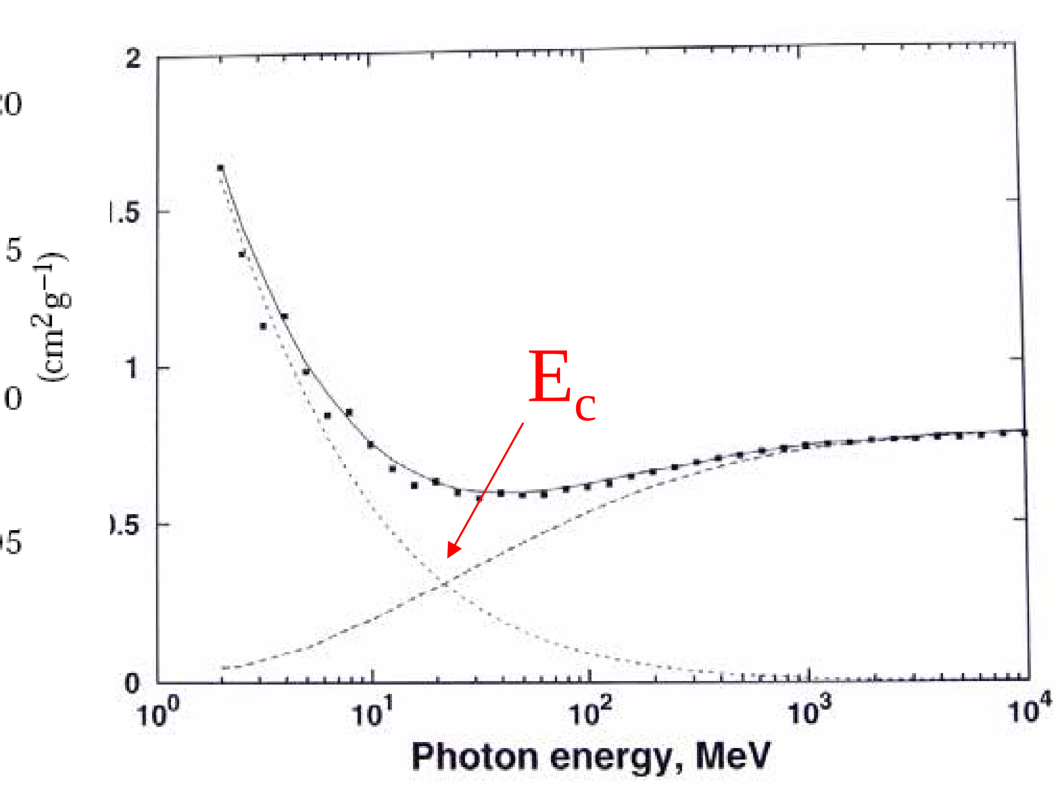
•energy loss $= 1/e$ due to bremsstrahlung

• 7/9 of the range of a γ due to pair production.

In air : $X_0 = 36.7 \mathrm{g/cm}^2$

EM cascades (Rossi & Greisen)

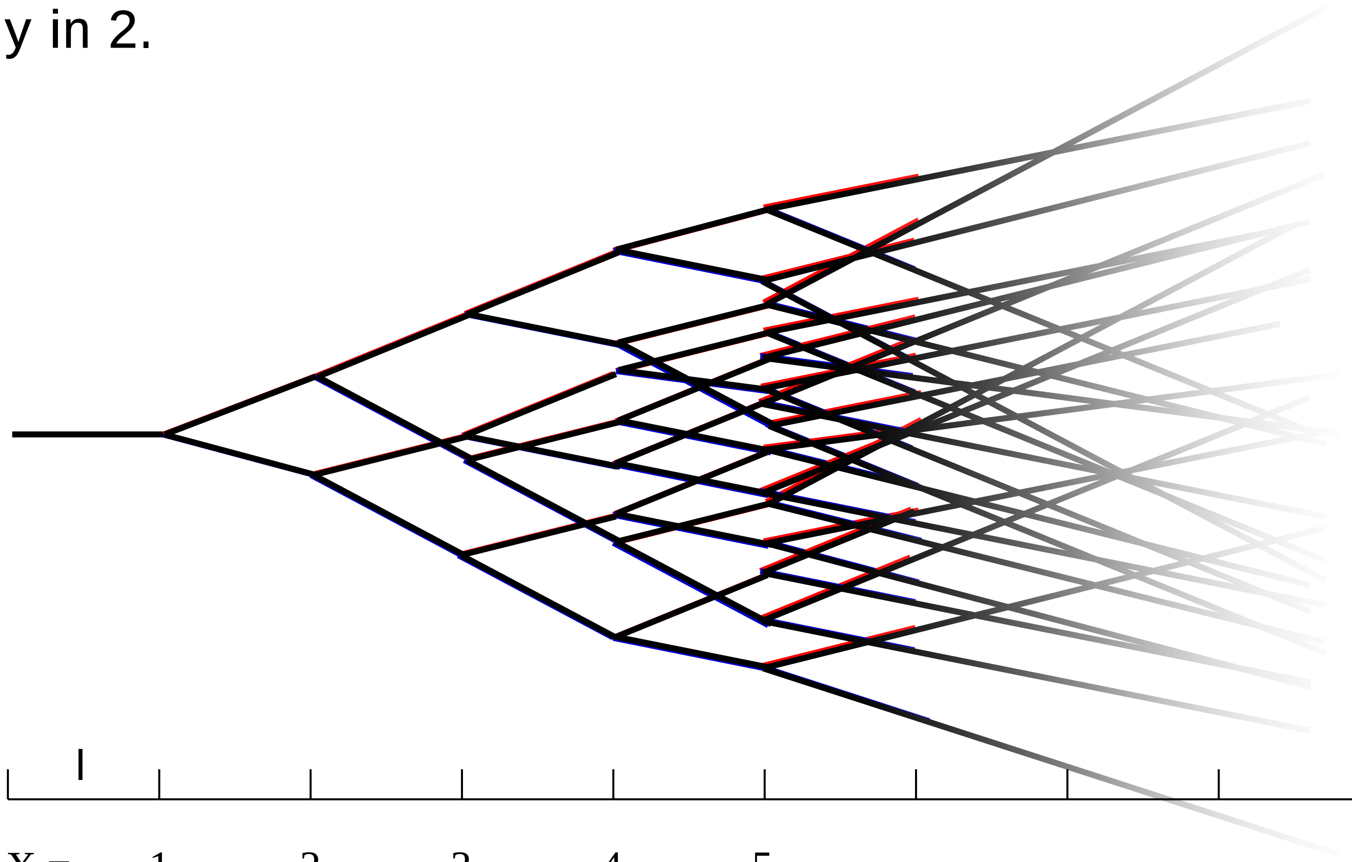




Critical energy: below this energy, ionization losses dominate.

Simplified development model (Heitler)

- Cascade consisting of only one type of particles having an interaction length l .
- At each interaction, 2 particles of same type are emitted sharing the energy exactly in 2.



X =	1	2	3	4	5
N =	2	4	8	16	32
E =	1/2	1/4	1/8	1/16	1/32

$E_c = E_0/32$

Longitudinal development

- After t radiation length, there are 2^t particles with energy

$$E = E_0/2^t$$
$$\text{soit : } t \ln 2 = \ln(E_0/E)$$

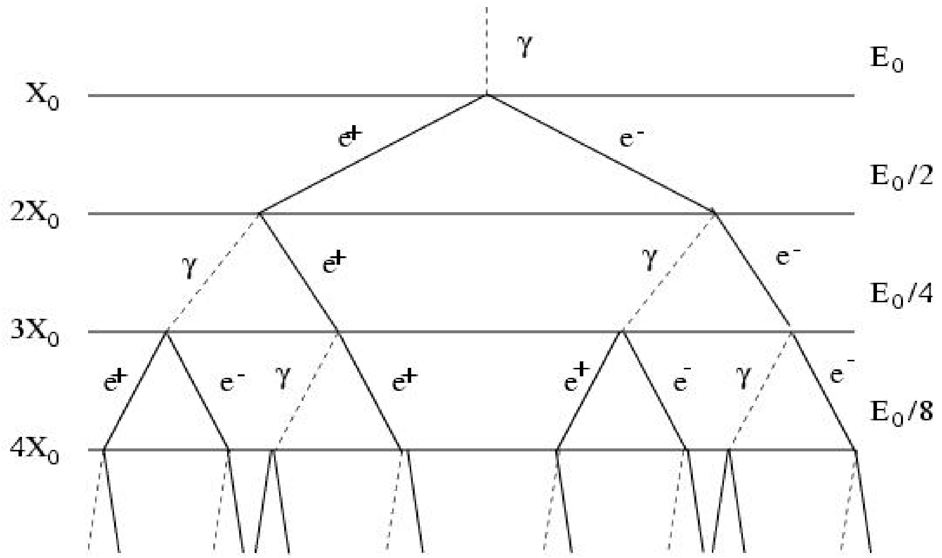
- The particles of energy E are produced at thickness:

$$t(E) \approx \ln(E_0/E)$$

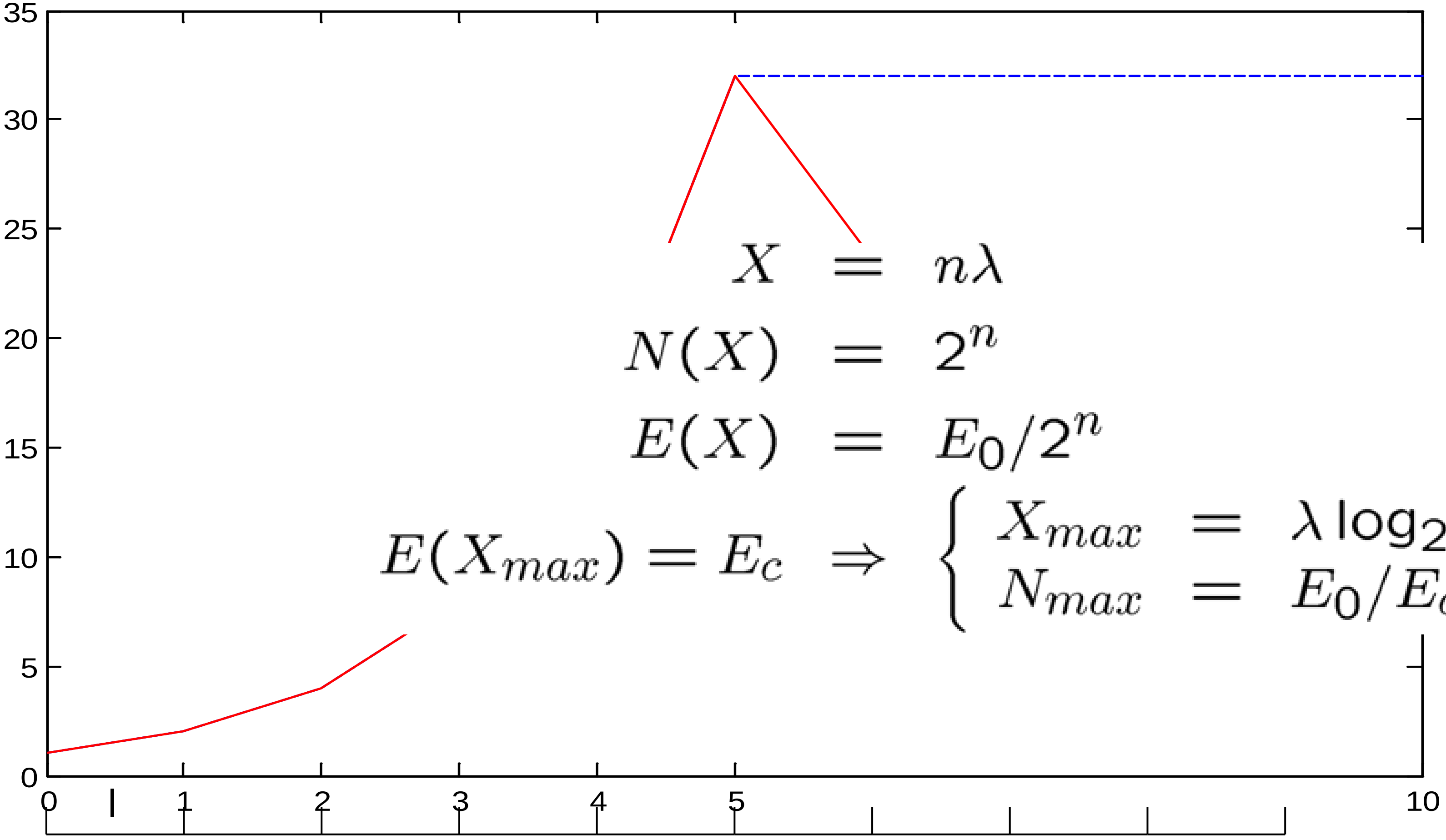
- The maximal development of the shower is reach for a thickness:

$$t_{max}(E_0) \approx ln(E_0/E_c)$$

- More realistic models agree with this rough estimate.



Simplified development model (Heitler)



$$
\begin{aligned}
X &= n\lambda \\
N(X) &= 2^n \\
E(X) &= E_0/2^n
\end{aligned}
$$

$$
E(X_{max}) = E_c \Rightarrow \begin{cases} X_{max} = \lambda \log_2(E_0/E_c) \\ N_{max} = E_0/E_c \end{cases}
$$

X =	1	2	3	4	5
N =	2	4	8	16	32
E =	1/2	1/4	1/8	1/16	1/32

$E_c = E_0/32$

Longitudinal development:
Approximation "A" (B. Rossi, K. Greisen)

- Approximation "A" describes the shower development phase where only bremsstrahlung and pair creation are in action.

- From Bethe-Heitler theory, one obtains des integro-differential linear and coupled equations leading to:

  - $\Pi(E,t)dE$ = average number of $e^{\pm}$ with energy $\in [E, E+dE]$, at tX_0 depth
  - $\Gamma(W,t)dW$ = average number of γ with energy $\in [W, W+dW]$, at tX_0 depth

- The simplifying factor is the absence of any energy scale.

Approximation A (cont)

- $\Pi(E,t)dE$ = average number of $e^{\pm}$ with energy $\in [E, E+dE]$, at tX_0 depth

- $\Gamma(W,t)dW$ = average number of γ with energy $\in [W, W+dW]$, at tX_0 depth

- Initial condition :

  – If the primary particle is a γ : $\Gamma(W,0) = \delta(E - E_0)$

  – If the primary particle is an $e^{\pm}$: $\Pi(E,0) = \delta(E - E_0)$

- Obvious special solutions:
  $\Gamma(W,t) = f(t)/W^{s+1}$ et $\Pi(E,t) = g(t)/E^{s+1}$
  (absence of energy scale)

... but they dont satify the initial conditions!

Approximation A (suite)

- The obvious solutions (power-law spectra, therefore scale invariant) correspond to an initial condition interesting in itself: an incident beam with a power law spectrum with an integral spectral index s.

- These special solutions form a base and a solution that fulfills the initial condition (photon or electron with an energy E_0) is obtained from a superposition of $1/E^{s+1}$ spectra (Mellin transformation, analogue to Fourier or Laplace transforms).

- Result : for a given value of t, the particle spectrum is very close to a power law $1/E^{s+1}$ with a value of s that varies with t and $y = \ln(E_0/E)$ following:

$$s = \frac{3t-1}{t+2y}$$

- The number of particle with energy E is maximal for $s = 1$

Taking into account ionisation energy losses: the "age" parameter

- Approximation A is not valid anymore when the electron mean energy energie is close to the cricital energy E_c.

- One can modify the above results:

$$y = \ln\left(\frac{E_0}{E_c}\right) \quad \text{et} \quad s = \frac{3t}{t+2y}$$

- Semi empirical formula given by Greisen for an incident γ, for the mean number of electrons after traversing t radiation length:

$$\bar{N}_t = \frac{0.31}{\sqrt{y}} \exp\left[t\left(1 - \frac{3}{2}\ln s\right)\right]$$

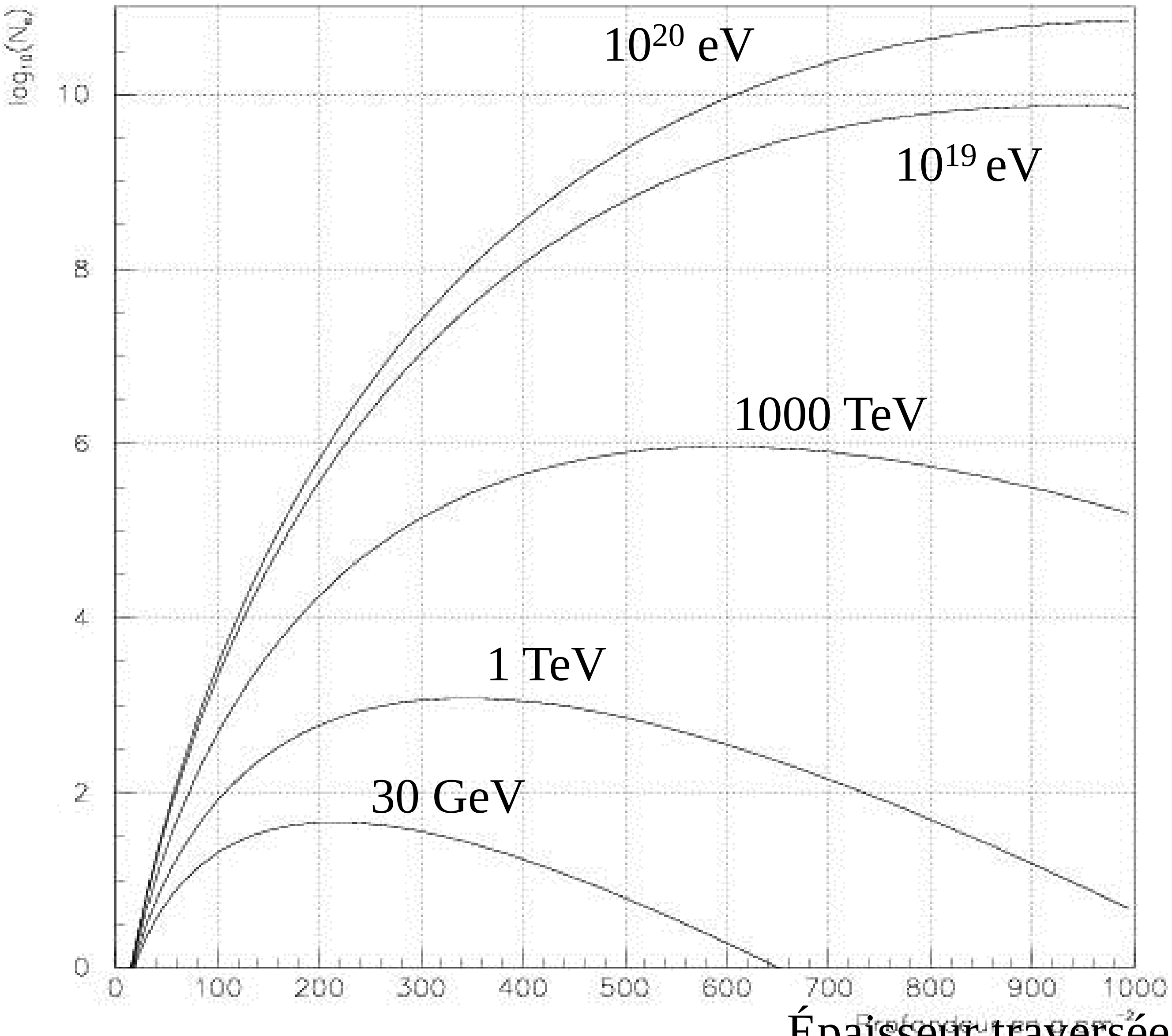
- The parameter s increase with t. It is < 1 during the development phase, reaches 1 at the maximal develpment stage for $t_{max} = y = \ln(E_0/E_c)$ and is > 1 during the extinction phase.

- s is called the "age".

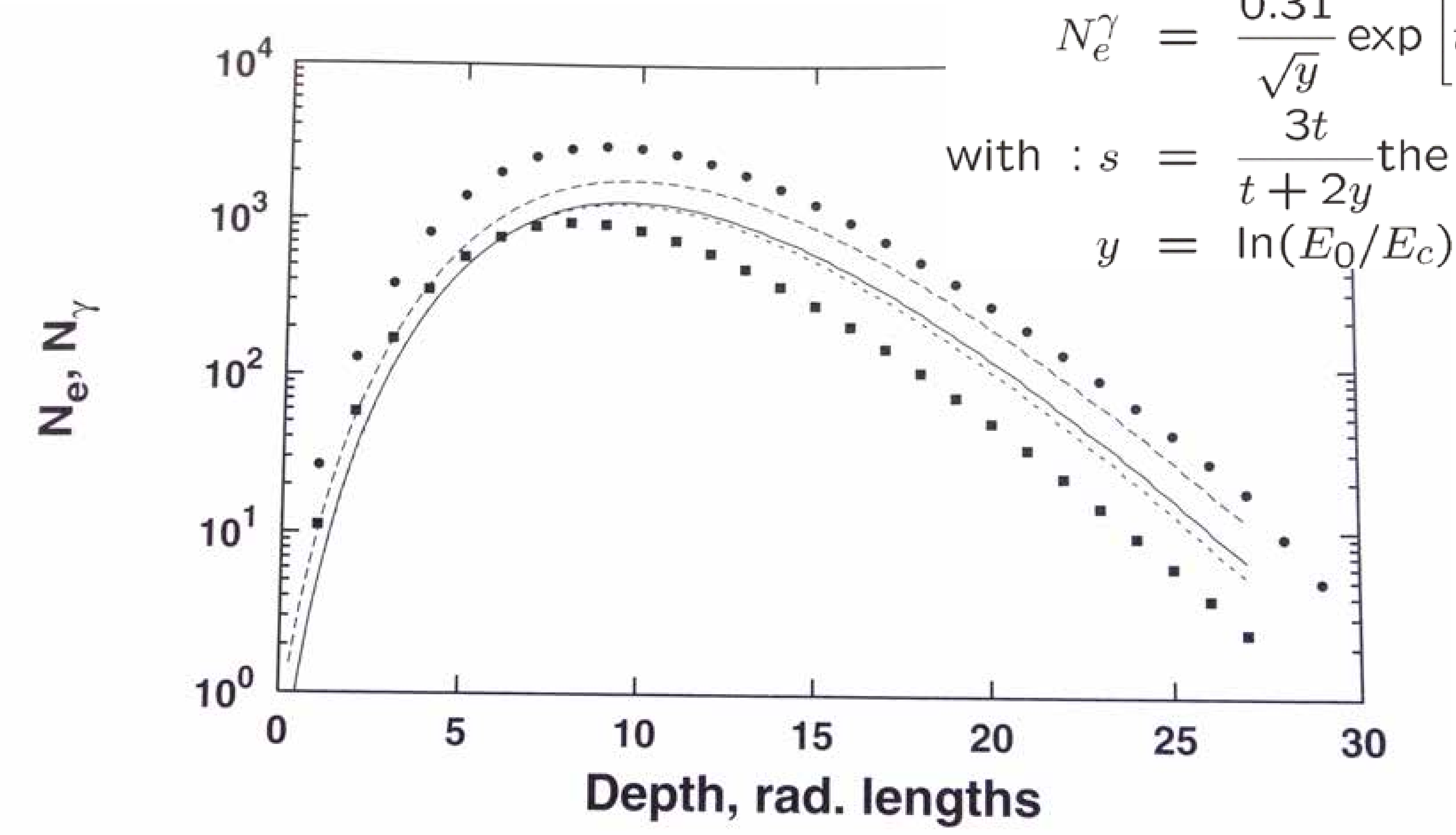
EM showers : some orders of magnitude

Primary genergy E_0	Thickness traverse $t_{max}\ X_0$ (g cm^{-2})	Altitude (m)	$N_e(t_{max})$
30 GeV	216	12000	50
1 TeV	345	8000	1200
1000 TeV	600	4400	$0{,}9 \times 10^6$
10^{19} eV	936	1200	$7{,}4 \times 10^9$
10^{20} eV	1021	0	$7{,}0 \times 10^{10}$

EM shower average profiles



EM cascades (Rossi & Greisen)



$$N_e^\gamma = \frac{0.31}{\sqrt{y}} \exp\left[t(1 - \frac{3}{2}\ln s)\right]$$

$$\text{with} : s = \frac{3t}{t+2y} \text{ the shower age}$$

$$y = \ln(E_0/E_c)$$

Parametrisation (Greisen) and Monte Carlo (EGS4)
photons 1TeV, E_c=10MeV

Shower size
i.e. number of electrons at ground level
as an energy estimator

- At maximal development level, the mean number of electrons quasi proportionnal to the primary energy ($y = \ln(E_0/E_c)$.

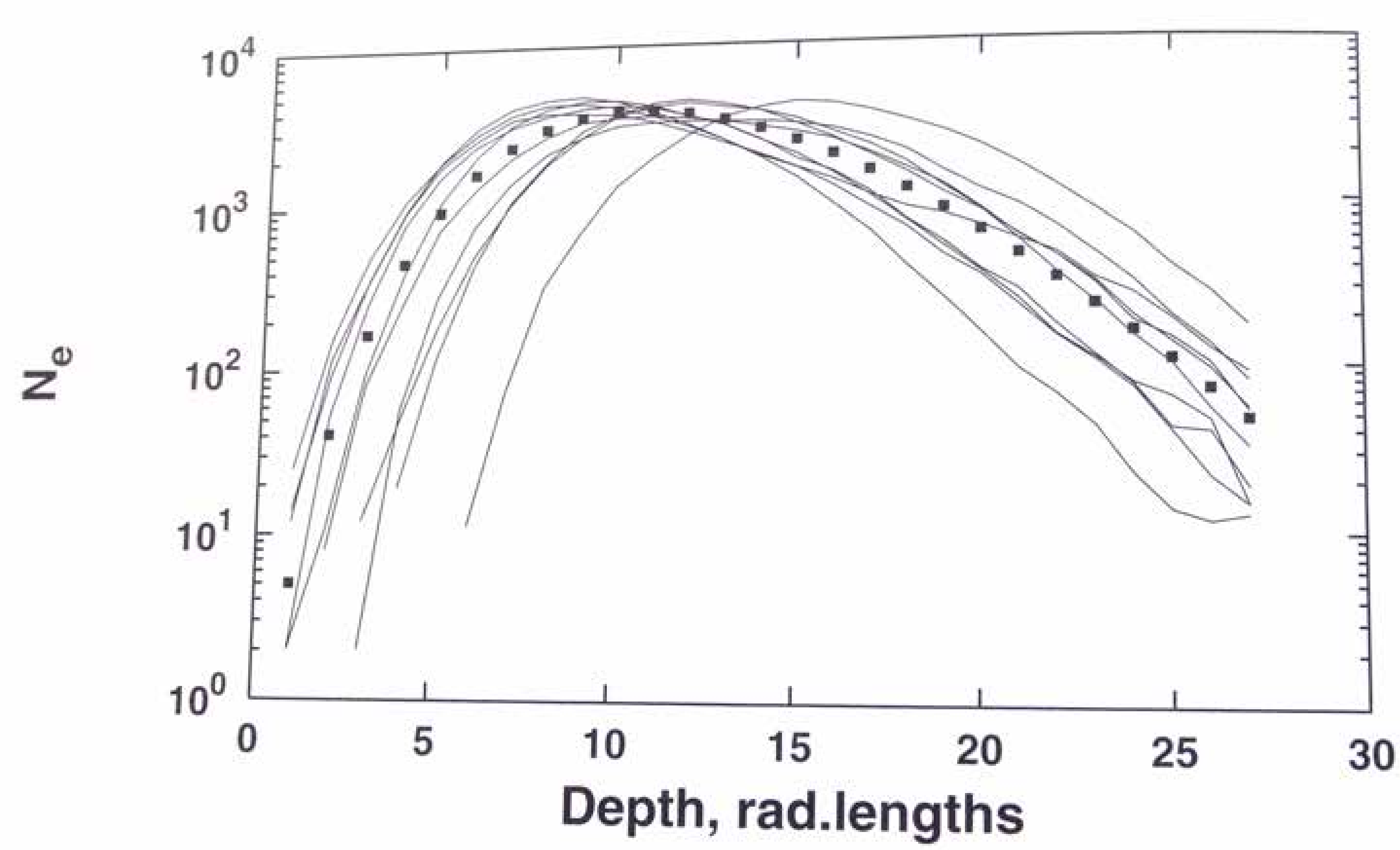
- Fluctuations on N_e:

  - Fluctuations on the depth of first interaction (exponential law)
  - Fluctuations in the shower development (approximatly log-normal because of the multiplicative behaviour)
  - Sampling fluctuations (depends on the type of detectors, their arrangement on the ground etc.)

- If the altitude of the maximal development is known (direct optical measurement), or if one can estimate the age independently (from lateral distribution of the electrons) one can avoid the first kind of fluctuations.

- Fluctuations are minimal at the maximum of development.

Cascades EM (Rossi & Greisen)

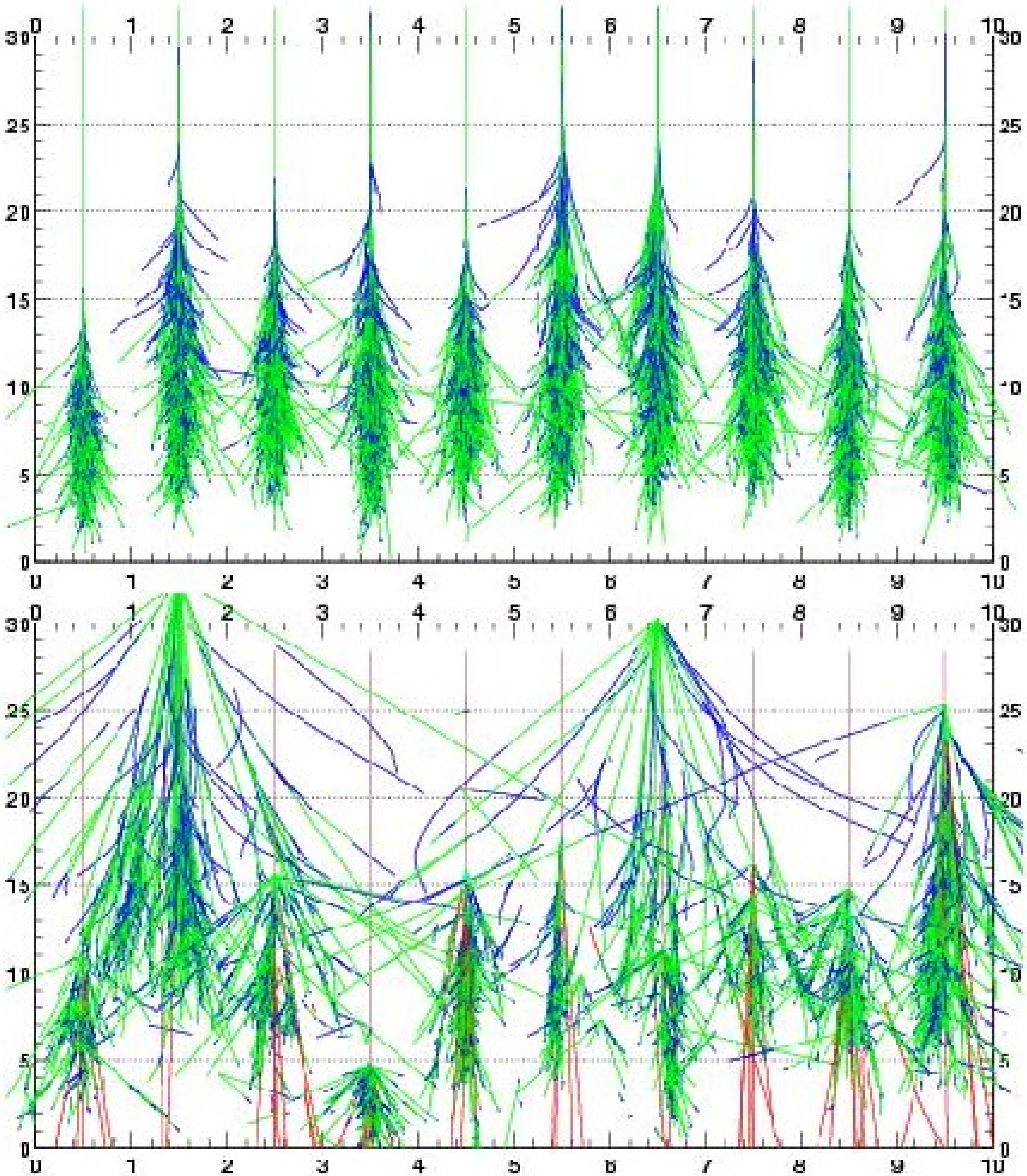


Shower to shower fluctuations
10 showers at 10^{14}eV compared to the average of 100 showers.

GAMMA-RAY (EM) INDUCED SHOWERS

10 γ
300 GeV

10 protons
300 GeV



Simulations de

M. de Naurois

Electromagnetic showers
($e^{\pm}$ or γ primary)

Dominating phenomena

- Radiation processes:
  - Bremsstrahlung of $e^{\pm}$
  - Pair production (>MeV) pairs $e^{+}e^{-}$
- Multiple scattering (small angular deflexions) of $e^{\pm}$

} In the coulombian field of nuclei

- Energy losses by $e^{\pm}$
  - par ionisation
  - excitation des atomes

γ induced shower 300 GeV

Roughly symmetric around the axis

Small transverse dispersion (multiple scattering)

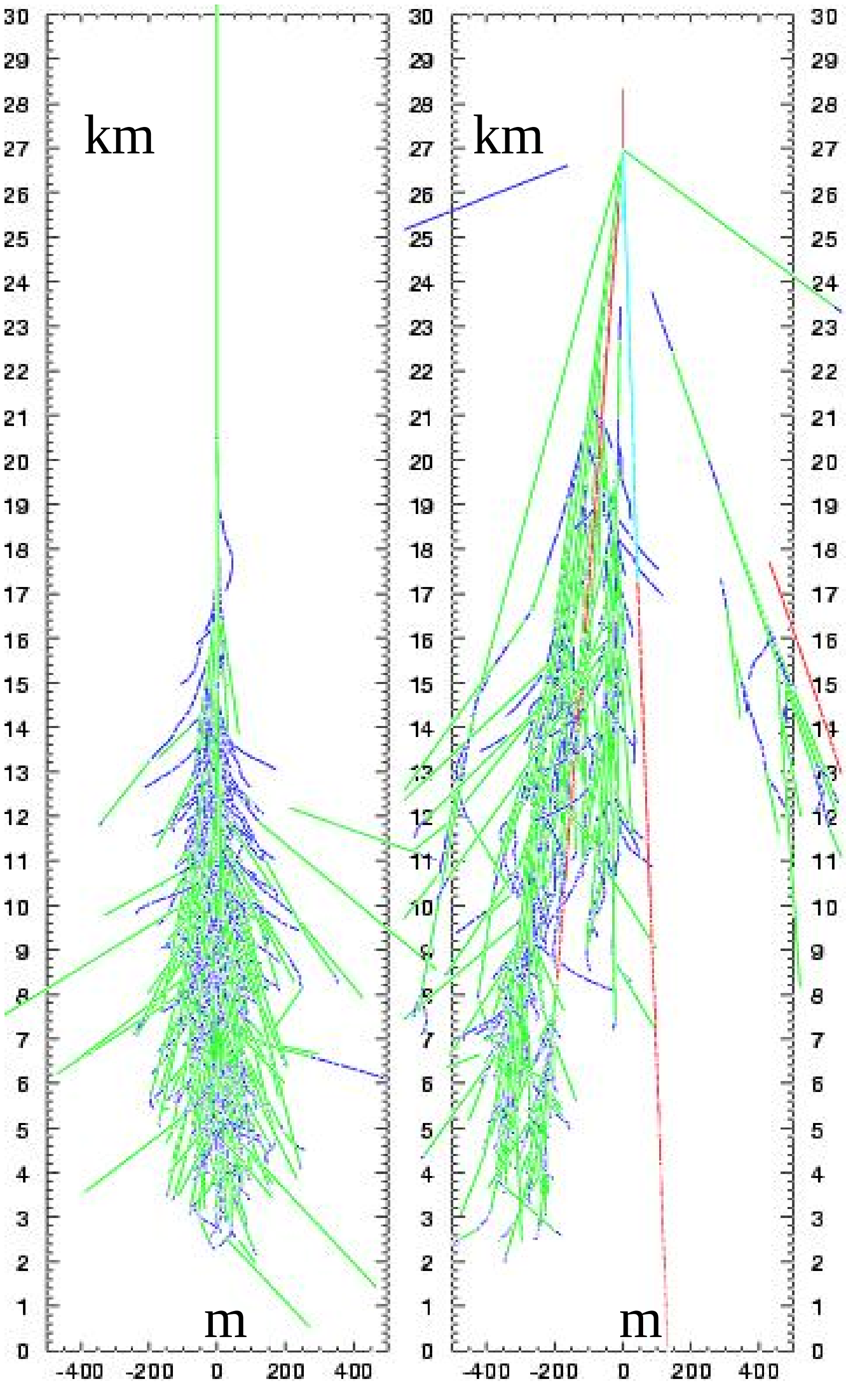
(almost) no muons

…

(unless E_0>1 PeV)

Essentially

e+ e- and γ secondaries



proton induced shower 300 GeV

Large transverse momentum

Muon component (from mesons decays)

A hadronic shower does contain EM sub-showers

Optical photon emission by showers

- Showers charged particles emit light:
  - Cherenkov light : very collimated along the shower axis (Cherenkov angle at 1 Atm. ≈ 1°) threshold depending on the altitude : at ground 22 MeV for $e^{\pm}$ et 4.5 GeV for $\mu^{\pm}$

    (20 photons per m per $\beta \approx 1$ charged particle at 1 atm)

    Essentially used for gamma-ray astronomy
  - Nitrogen fluorescence: isotropic emission

    (¼ 4 photons per electron per m)

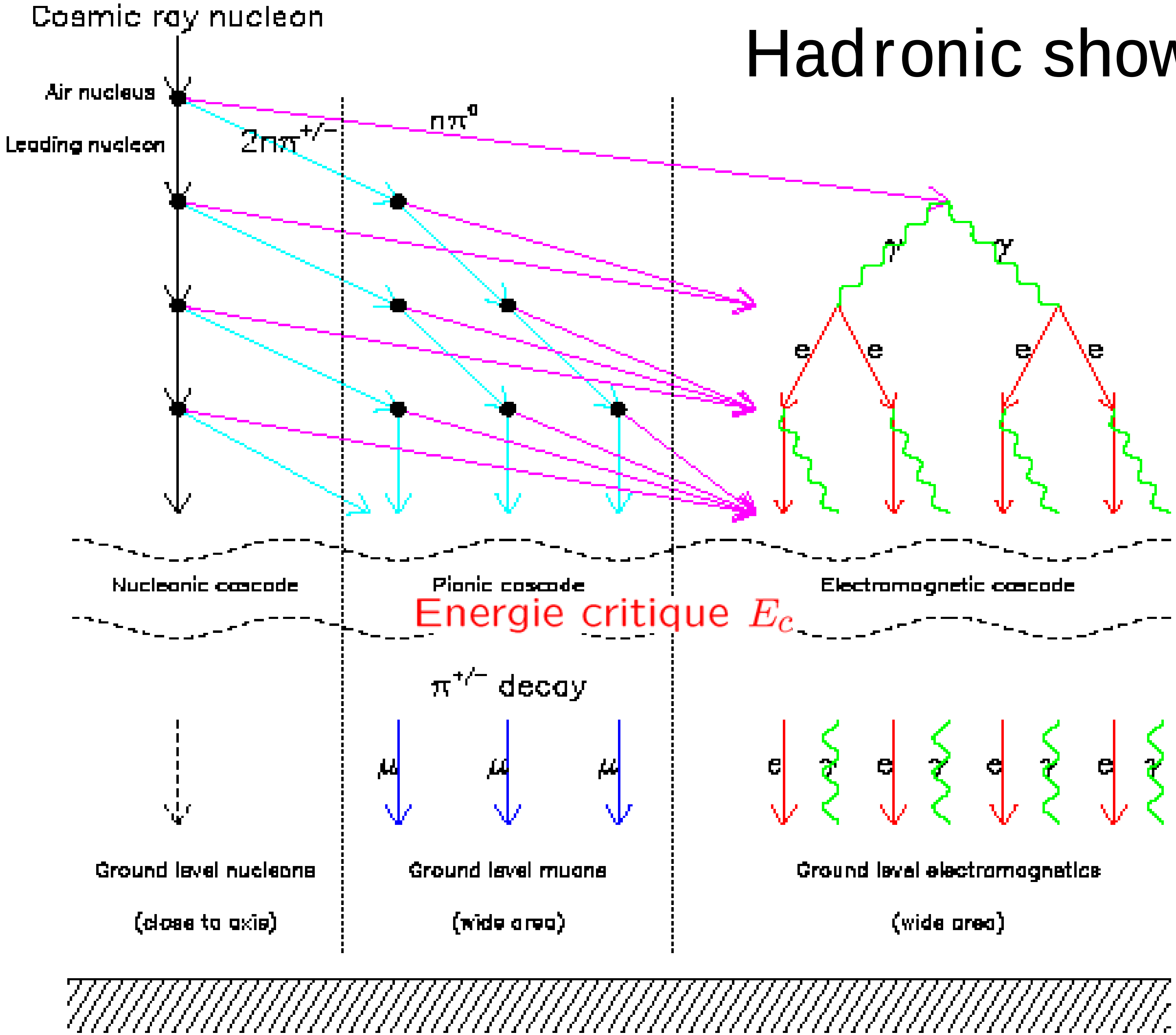
    Essentially used at UHE ≥ 10^{18}eV.

Lecture on Imaging & Cherenkov Detectors

- This light detected by ground telescopes gives us very rich information on the 3D development of the showers. It give a quasi calorimetric reliable measurement of the energy.
- ... but optical detectors can only work during moonless clear sky nights (≈ 10% duty cycle).

HADRONIC SHOWERS MODELS AND DETECTION

Hadronic showers



"Hadronic" showers
(proton ou noyau primaire)

- Great complexity implying the use of numerical simulations:
  - Many length scales : nucleon interaction length, pion interaction length, EM radiation length, atmosphere density height scale...
  - Superposition of a nuclear cascade, a pionic cascade and an electromagnetic cascade (the later from π^0 decay γ).
  - Large fluctuations in the multiplicity of secondaries.
- Bur simulations are subject to many uncertainties:
  - p+N or N+N interactions: sensitivity to nuclear models.
  - Energy range unexplored by accelerators and colliders : sensitivity to nucleon structure functions (parton distributions) and fragmentation functions extrapolated far from the measured regions.
  - The inelasticity and in general the very forward diffractive physics is not well measured in fixe target experiment (even worse at colliders).

  Still, the main behaviour observed on EM showers remains valid.

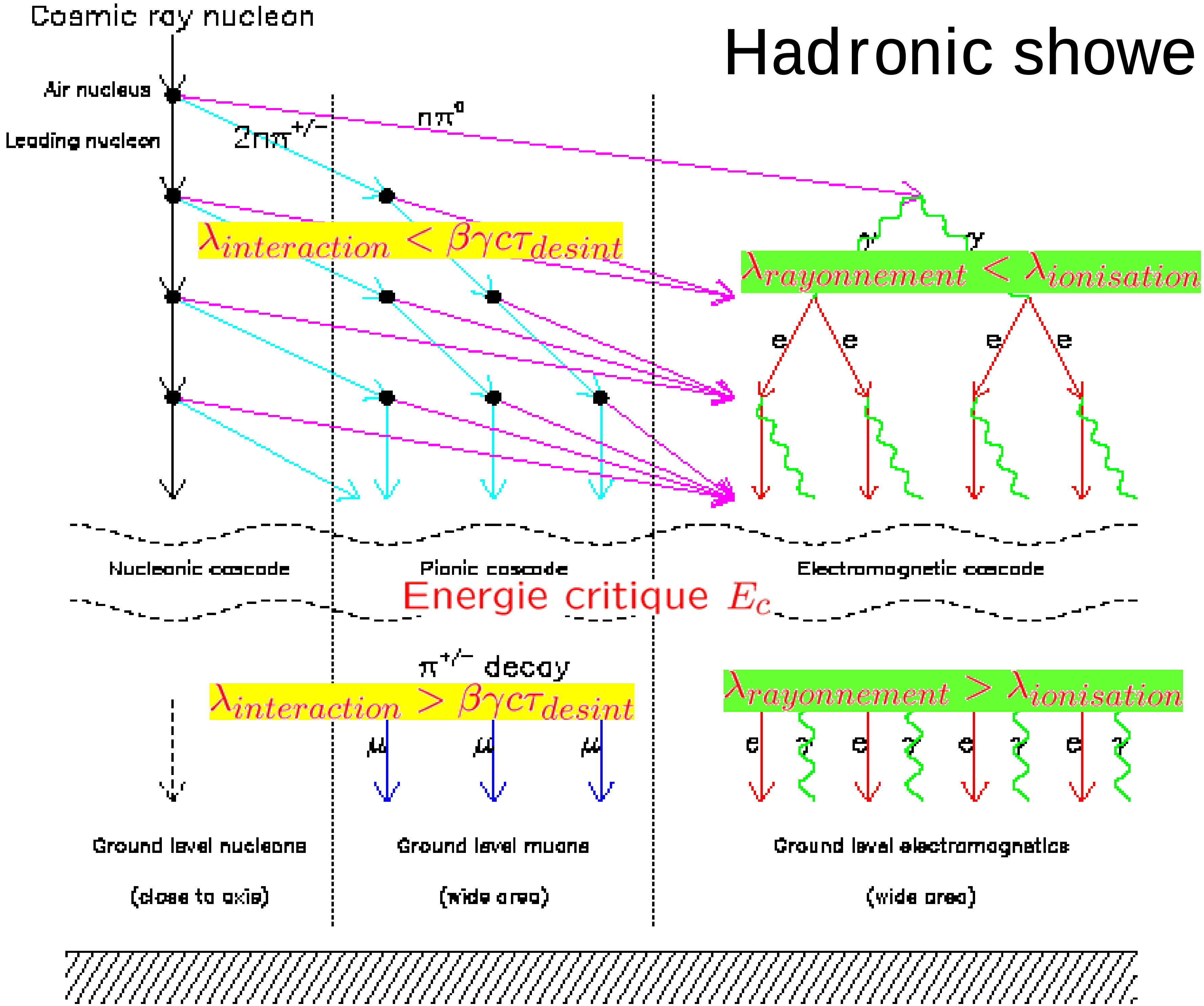
From EM to Hadronic showers

- The main observables are the same:
  - Number of electrons, gamma but also muons at ground and their lateral distributions.
  - Longitudinal profil and maximal dev. altitude (optical detectors).
  - Number of muons at ground level and lateral distribution of muons.

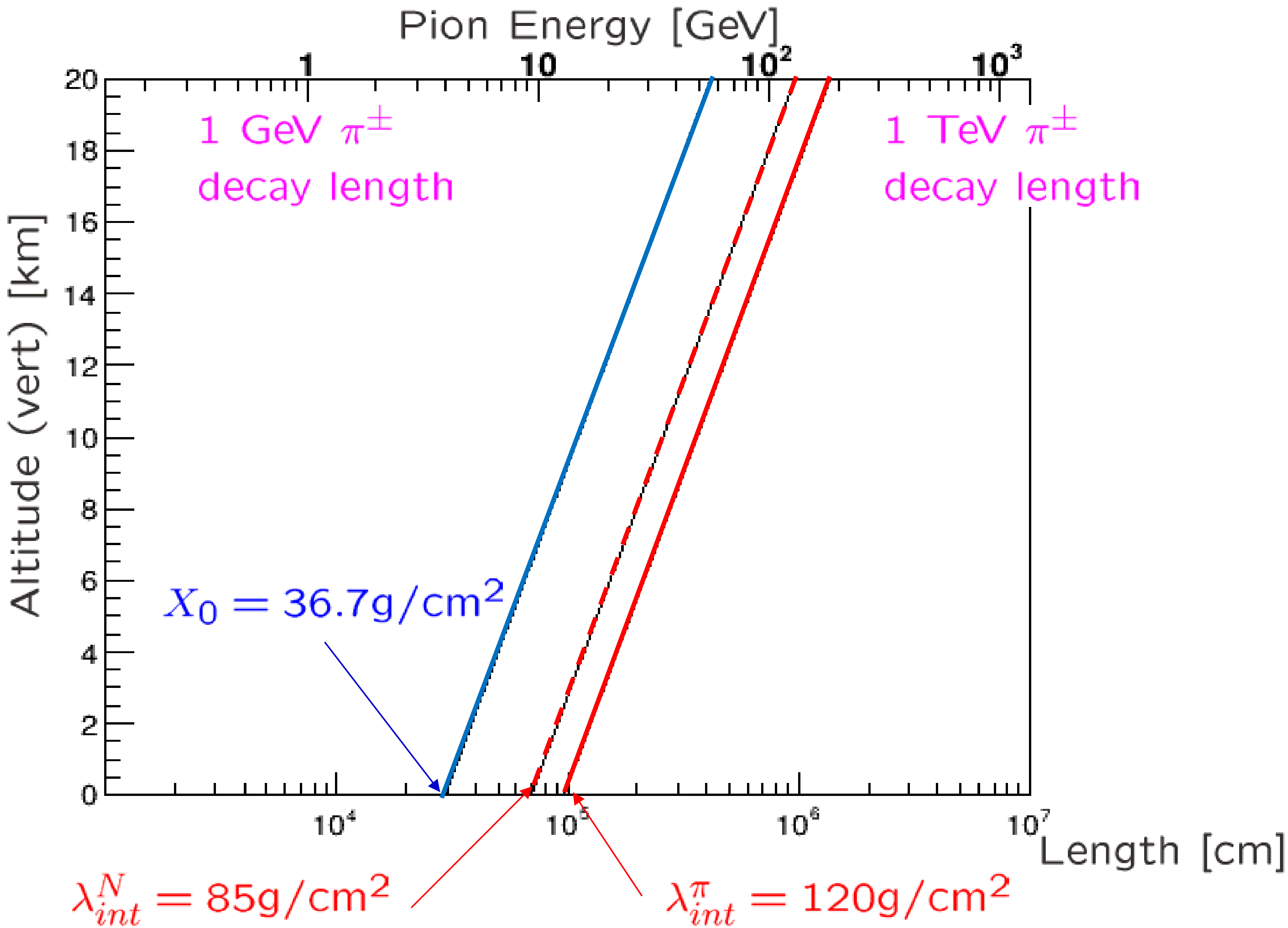
- Feynman scaling is rather well verified in the fragmentation: it plays an role analogue to that of Bethe-Bloch formulae for EM showers (absence of mass/energy scale).

- Simulations have allowed to establish empirical formulae inspired by EM showers useful to to quite estimates *(T.K. Gaisser, A.M. Hillas)*

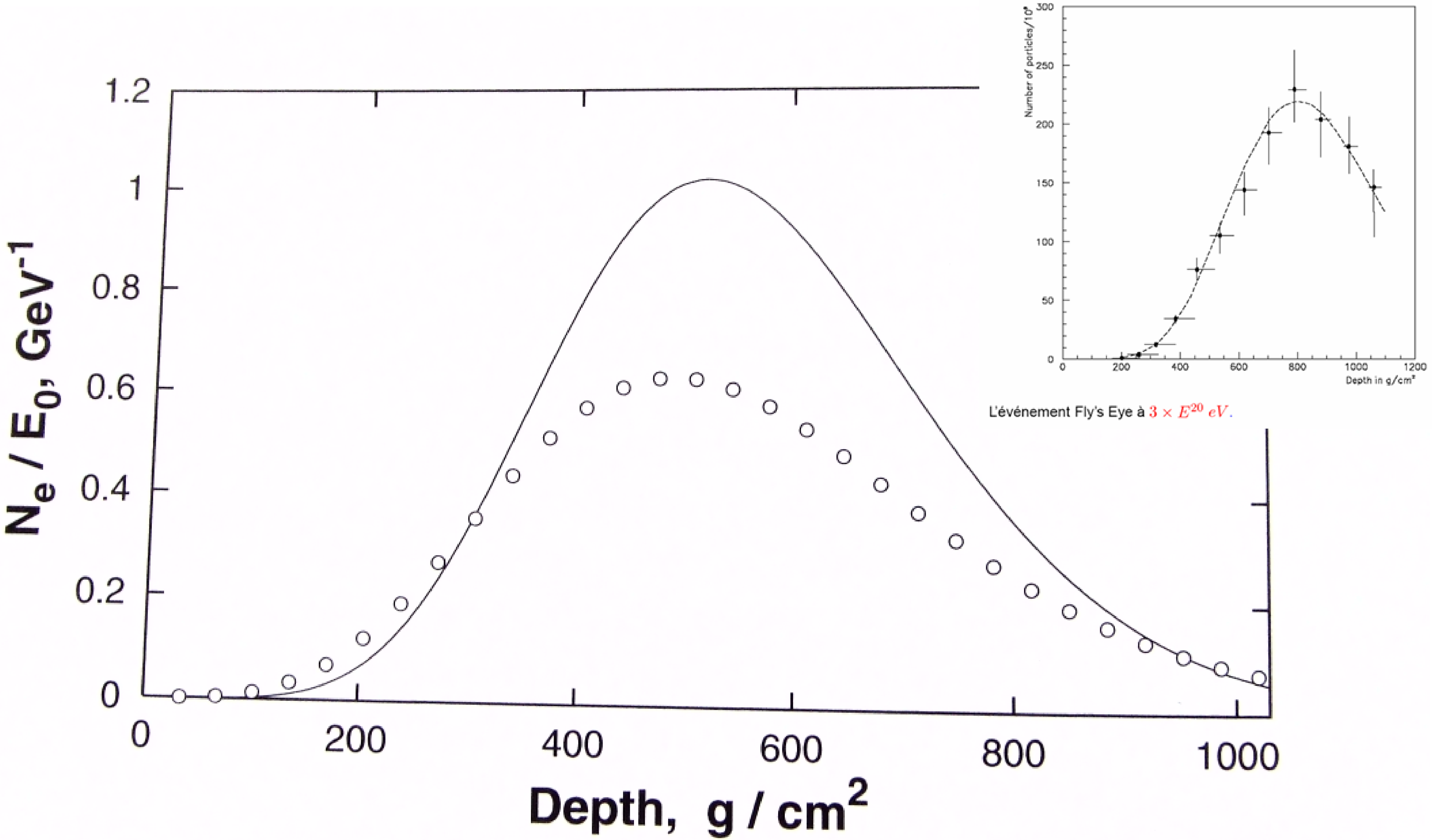
Hadronic showers



Interaction and radiation lengths in atmosphere



Development of Hadronic vs EM showers



L'événement Fly's Eye à $3 \times E^{20}\ eV$.

Gaisser longitudinal Parametrisation

Gaiser Hillas formulae :

$$N_e(X - X_1) = N_e^{max} e^p \left(\frac{X - X_1}{X_{max} - \lambda}\right)^p \exp - \left(\frac{X - X_1}{\lambda}\right)$$

$$\text{avec } p = \frac{X_{max} - \lambda}{\lambda}$$

Averaging on X_1 depth of 1^{st} interaction :

$$\bar{N}_e(X) = N_e^{max} \frac{p}{p+1} e^p \left(\frac{X}{X_{max} - \lambda}\right)^{p+1} \exp - \left(\frac{-X}{\lambda}\right)$$

$$X_{max} = X_0 \log\left(\frac{E_0}{\epsilon_0}\right)$$

$$N_e^{max} = \frac{E_0}{\omega}$$

Radiation length : $\approx 36.7 \text{g/cm}^2$

Cricital energy : $\epsilon_0 \approx 74\text{eV}$

Empirical relation between size and energy: $\omega \approx 1.7\text{GeV}$

Incident nucleus interaction length (of energy E_0) $\lambda_N \approx 70\text{g/cm}^2$

Longitudinal develpopment

Xmax and energie :

$$X_{max} \approx X_0 \log\left(\frac{E_0}{\epsilon_0}\right)$$
$$\Rightarrow 80 g/cm^2 \text{ per energie decade}$$

Nuclei :

Superposition principle : a nucleus ^{A}N is equivalent to A protons.

Thus :

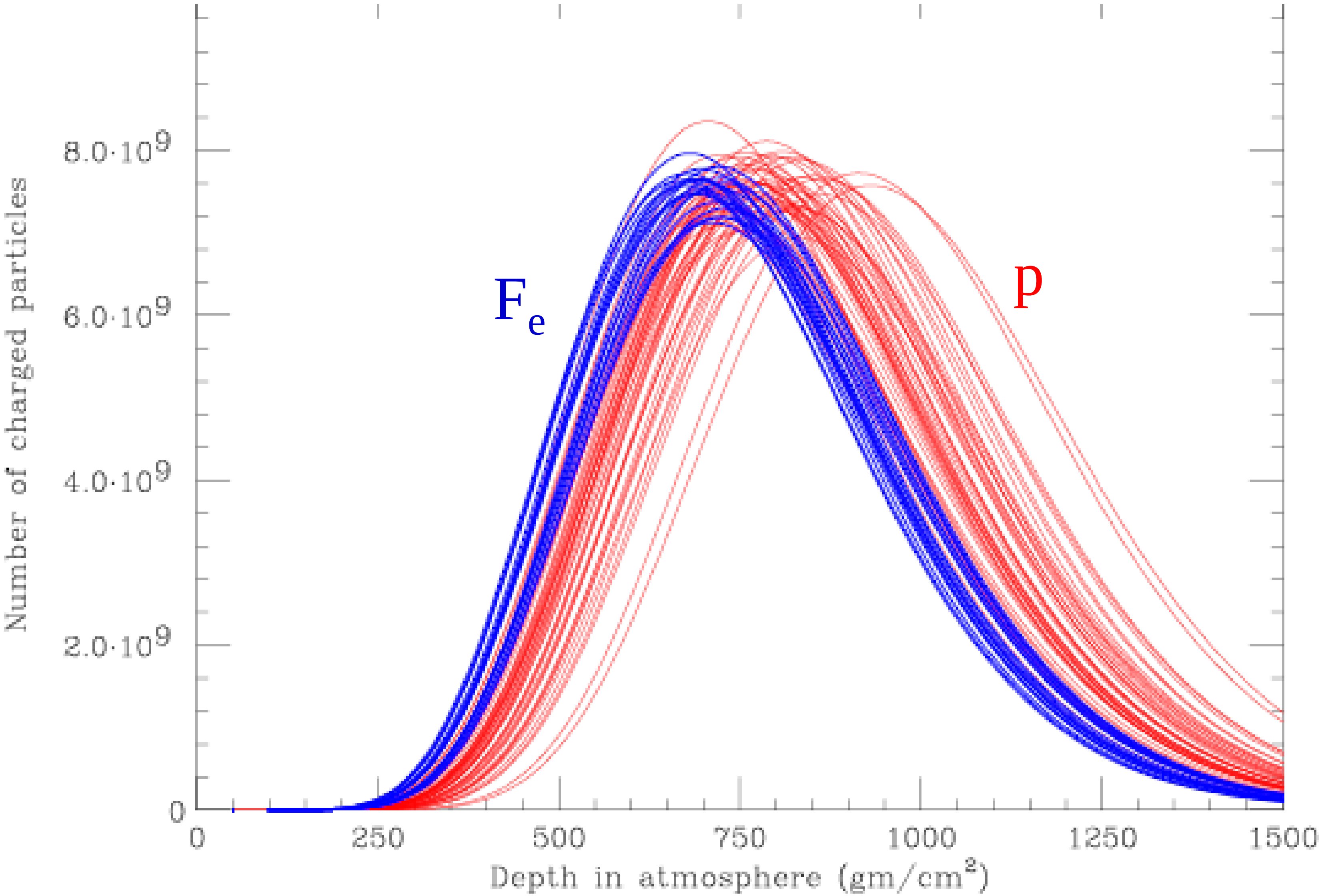
$$X_A^{max} = X_0 \log\left(\frac{E_0}{A\epsilon_0}\right)$$
$$= X_p^{max} - X_0 \log(A)$$

For example iron/proton $A = 56$:

$$X_0 \log(A) = 36.7 \log(56) = 148 g/cm^2$$

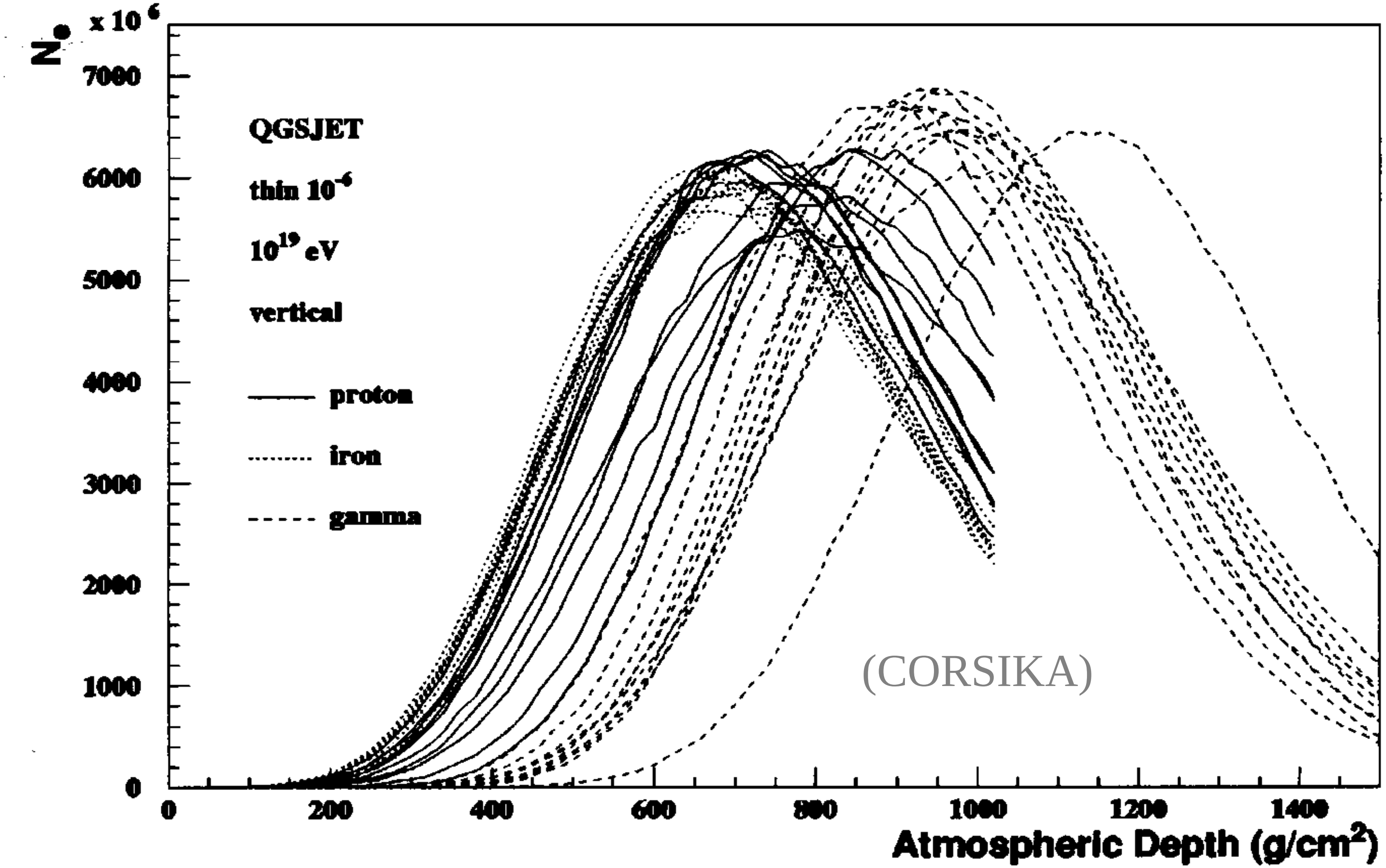
Structure in space

Shower to shower fluctuations largely due to the depth of the first interaction.



Primary identification

- Requires a good statistics and a good knowledge of the initial energy, the shower angle (+ systematic corrections because of atmospheric attenuation)



Radial extention

The radial distribution is determined by the mean transverse momentum (P_T) from hadronic interactions and by multiple scattering. In air, the Molière radius is $\approx$ 75m.

Molière radius (~1/4 of the radiation length) :

$$\left\langle \delta\theta^2 \right\rangle = \left(\frac{21\text{MeV}}{E}\right) \delta X$$

$$r_1 = \left(\frac{E_s}{E_c}\right) X \approx 9.3 \text{ g/cm}^2$$

Nishimura, Kamata, Greisen :

multiple scattering + transverse momentum

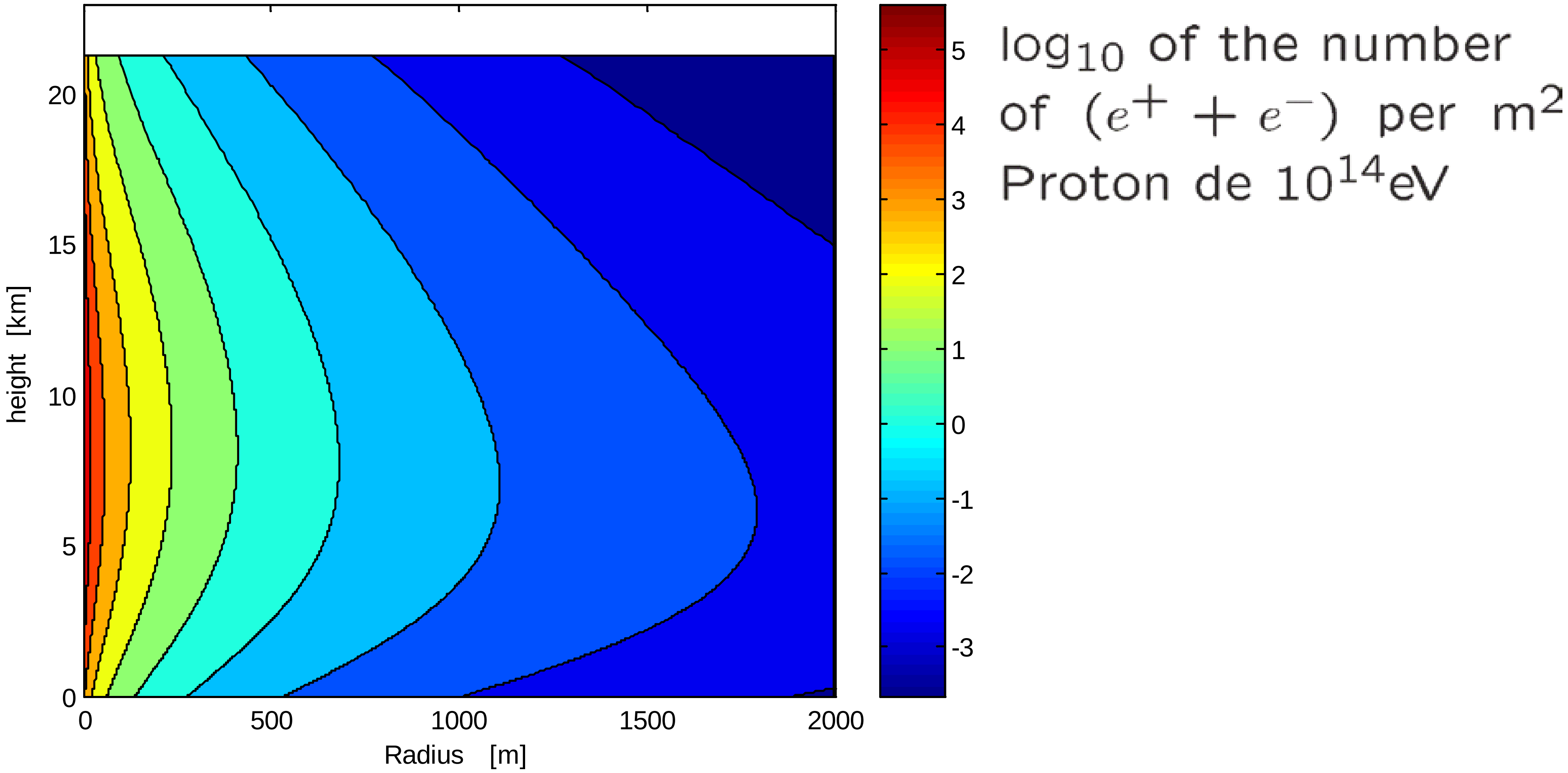
$$xf(x) = C(s)x^{(s-1)}(1+x)^{(s-4.5)}$$

$$\text{with } : x = \frac{r}{r_1}$$

normalization tel que :

$$2\pi \int_0^\infty xf(x)dx = 1$$

$e^+ + e^-$ lateral density



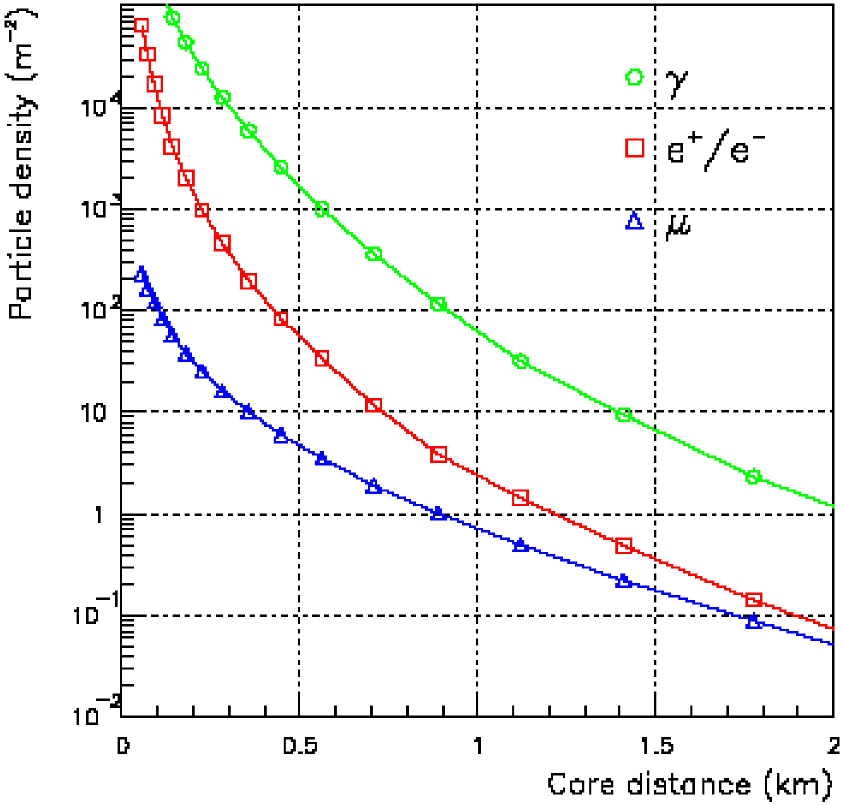
Lateral evolution

The density as a function of the distance to the center of the shower is characterized by a lateral desity function (LDF)

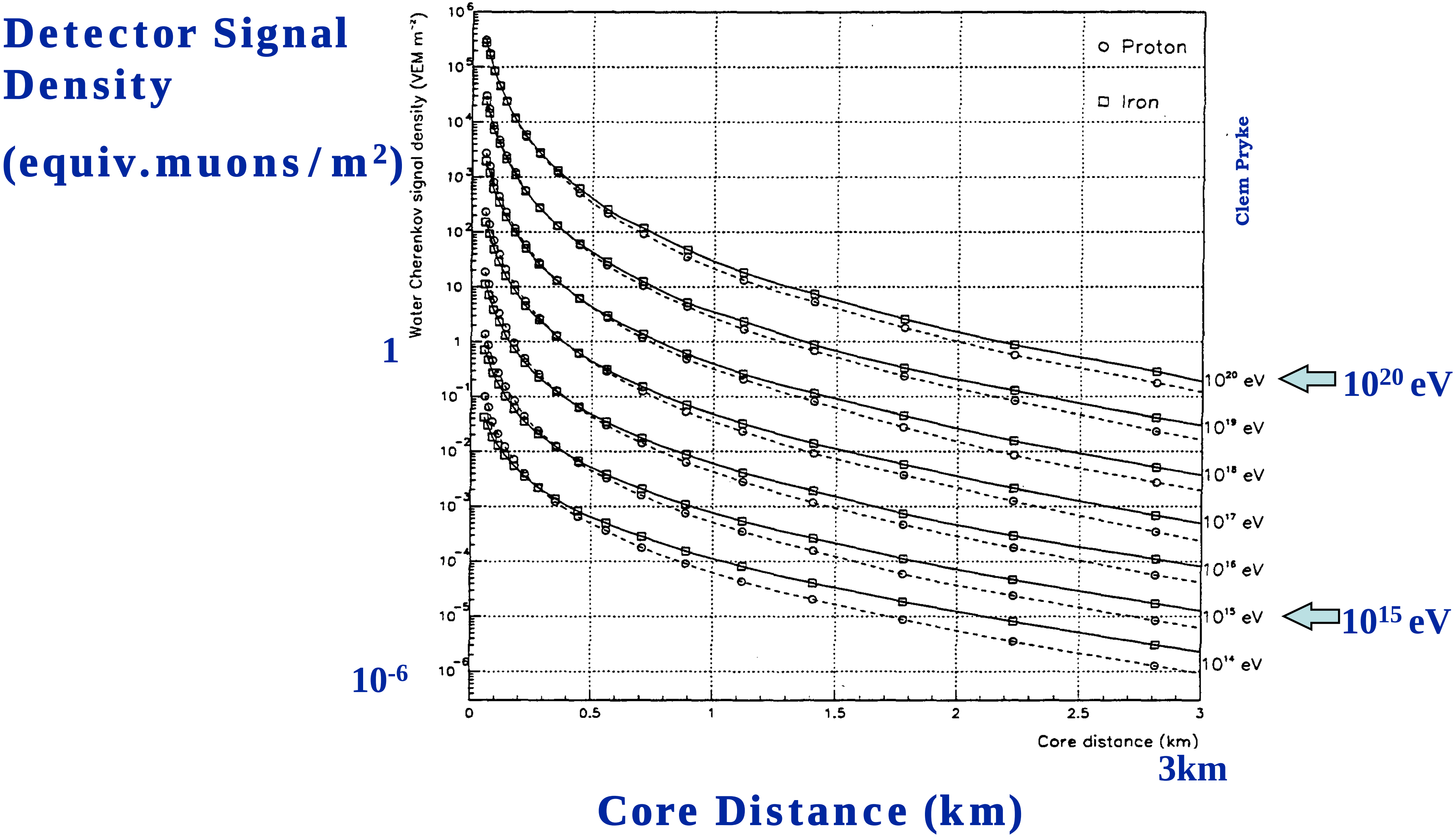
$$\rho(r) \propto k \times r^{-[\eta+f(r)]}$$

where f et k depends on the type of detectors used where η depends on the incident angle of the shower and the primary energy.

For $r > 800$m this (empirical) expression must be modify as $(r/800)^{1.03}$



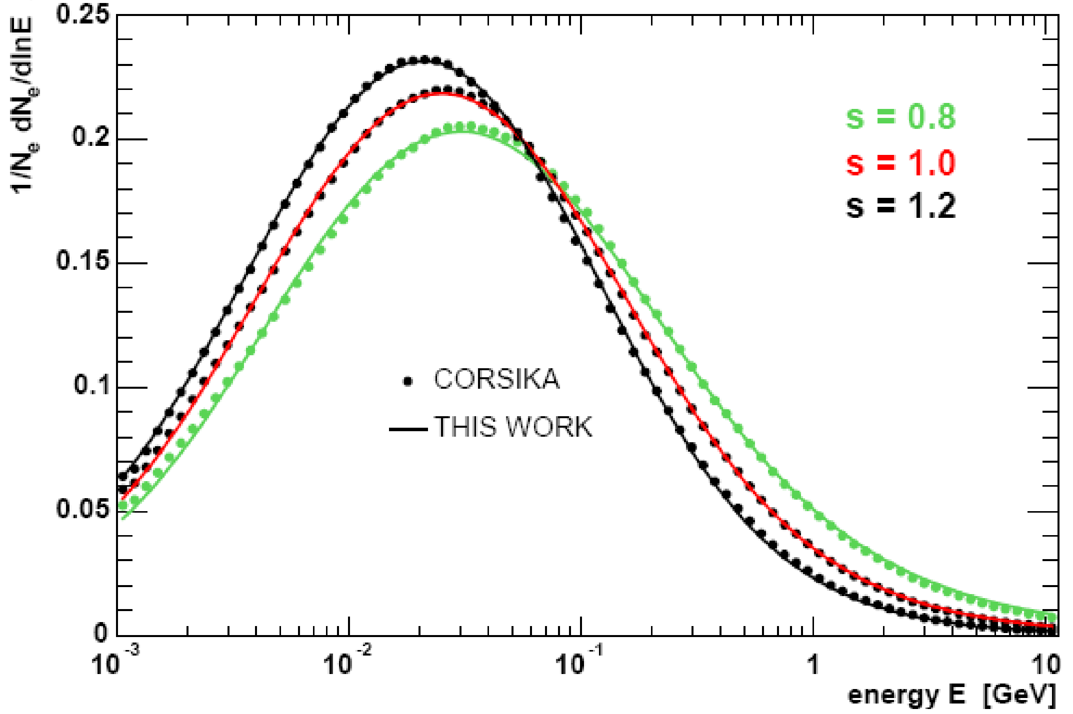
Shower Density Lateral Distribution
(simulation)



Particle energy distribution

Rossi Greisen :

$$\frac{dN}{d(\log E)} \approx \frac{1}{E^{s+1}}$$

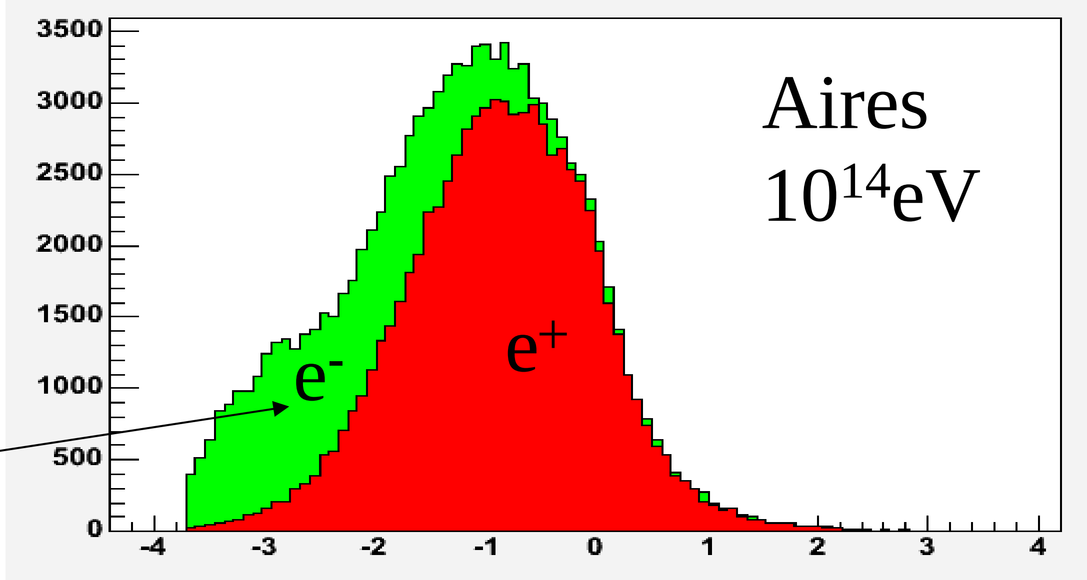


E.Nerling (thesis) :

$$\begin{aligned} f_e(E,s) &= a_0 \frac{E}{(E+a_1)(E+a_2)^s} \\ a_1 &= 6.42522 - 1.53183.s \\ a_2 &= 168.168 - 42.1368.s \end{aligned}$$

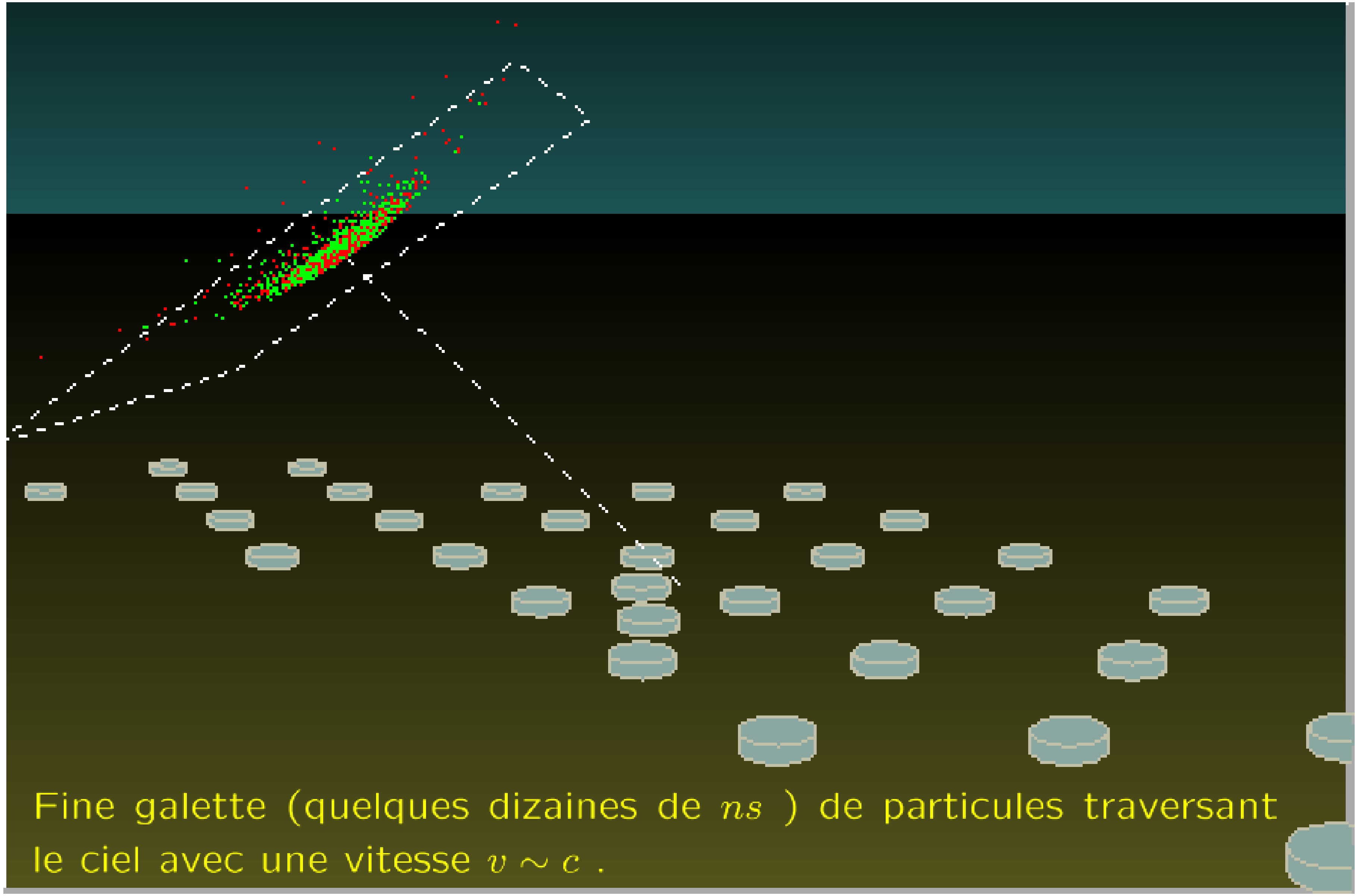
E en MeV

$$a_0^{-1} = \int_{\log E_{cut}} f(E,s)d(\log E)$$

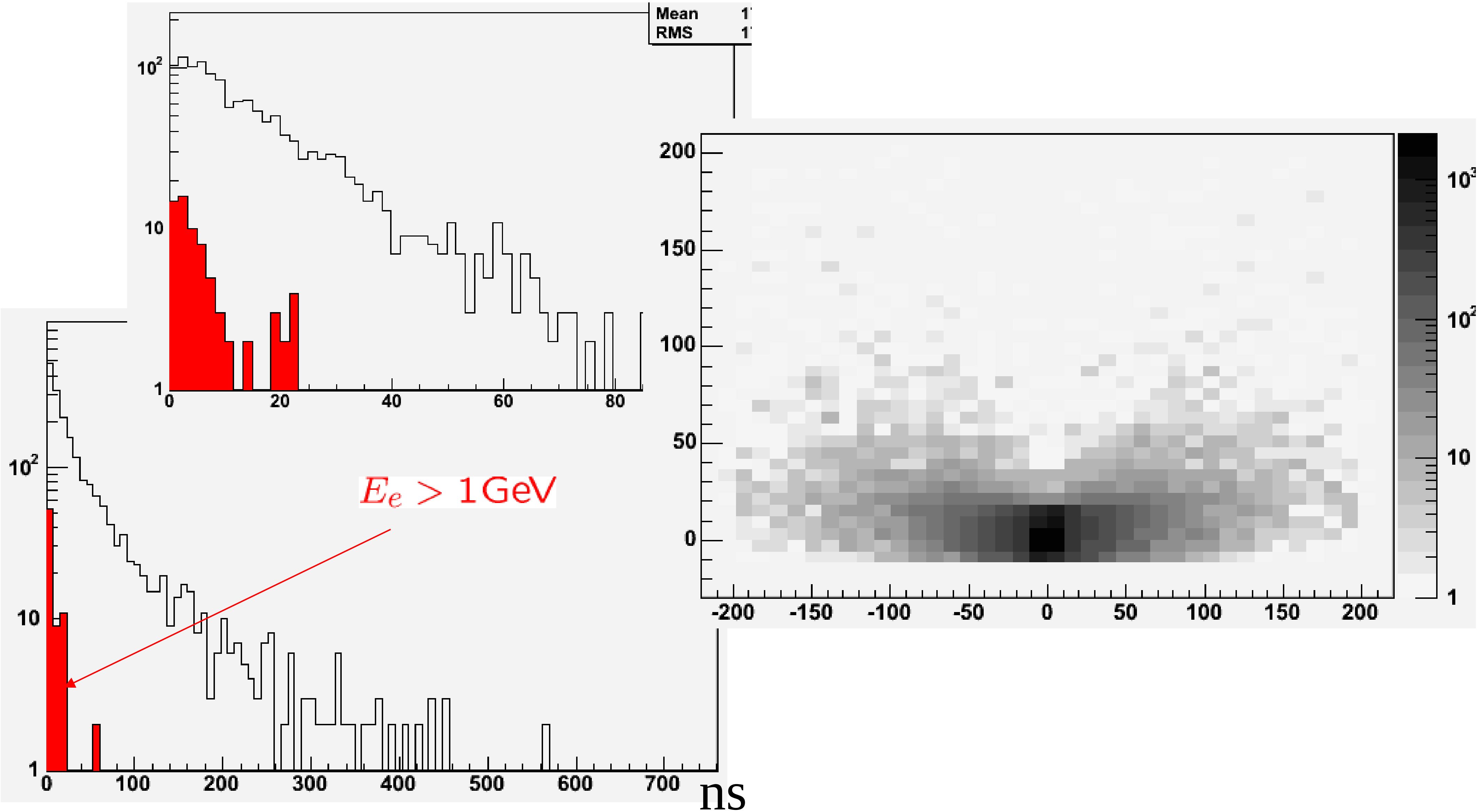


Excess e- at low energy (ionization)

Time structure



Time structure



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UHECR detection

Lecture on
Imaging & Cherenkov
Detectors

Neutrino Physics with astroparticules

Lecture on
Imaging & Cherenkov
Detectors