

SSI<sub>2014</sub>

# Shining Light on **DARK MATTER**

Structure Formation

Tom Abel

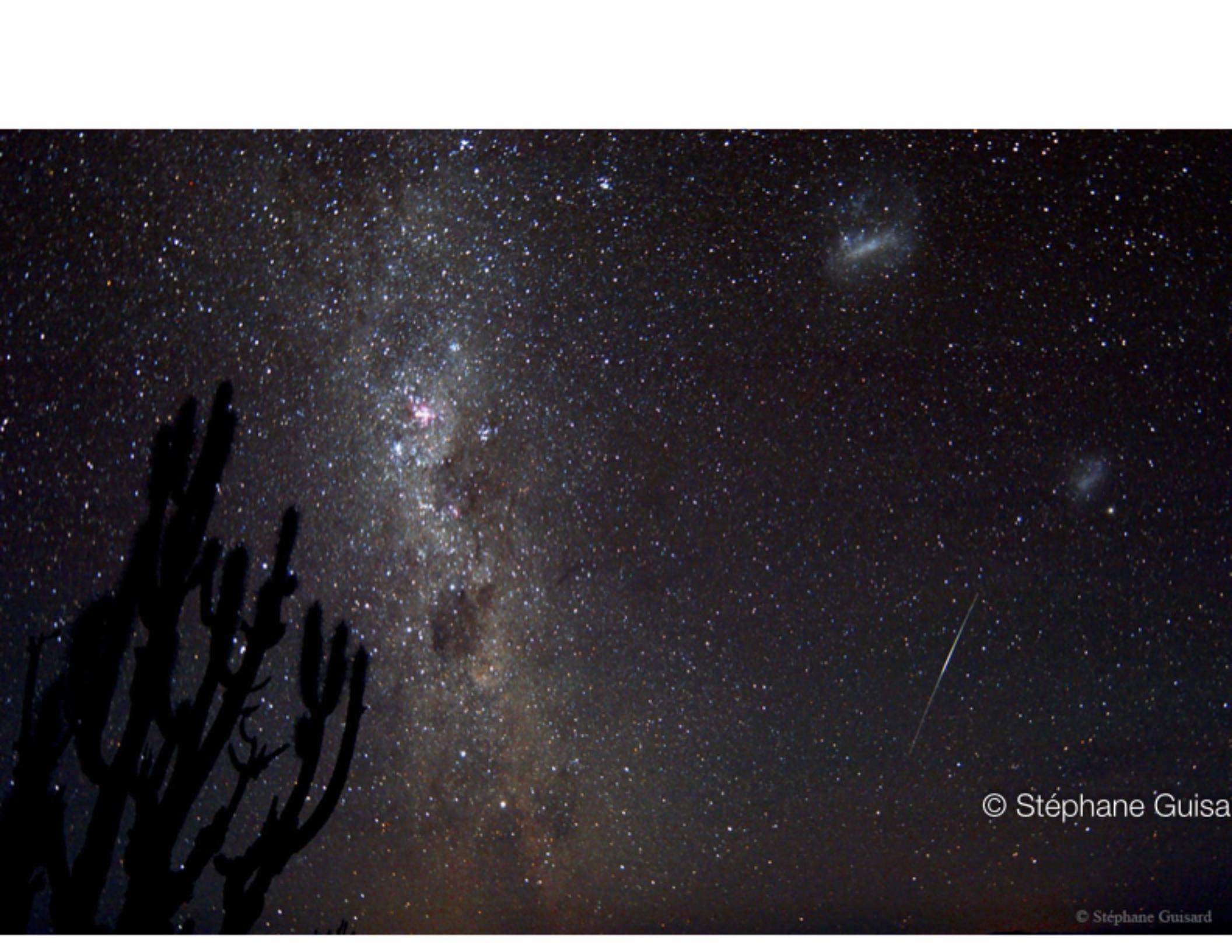
# Outline

- What Structure?
- Gravitational instability forms structure in the Universe
- Vlasov-Poisson and N-body
- Numerical Techniques & Codes
- Cold, warm & hot Dark Matter models
- DM Halo properties and other select Simulation results



today

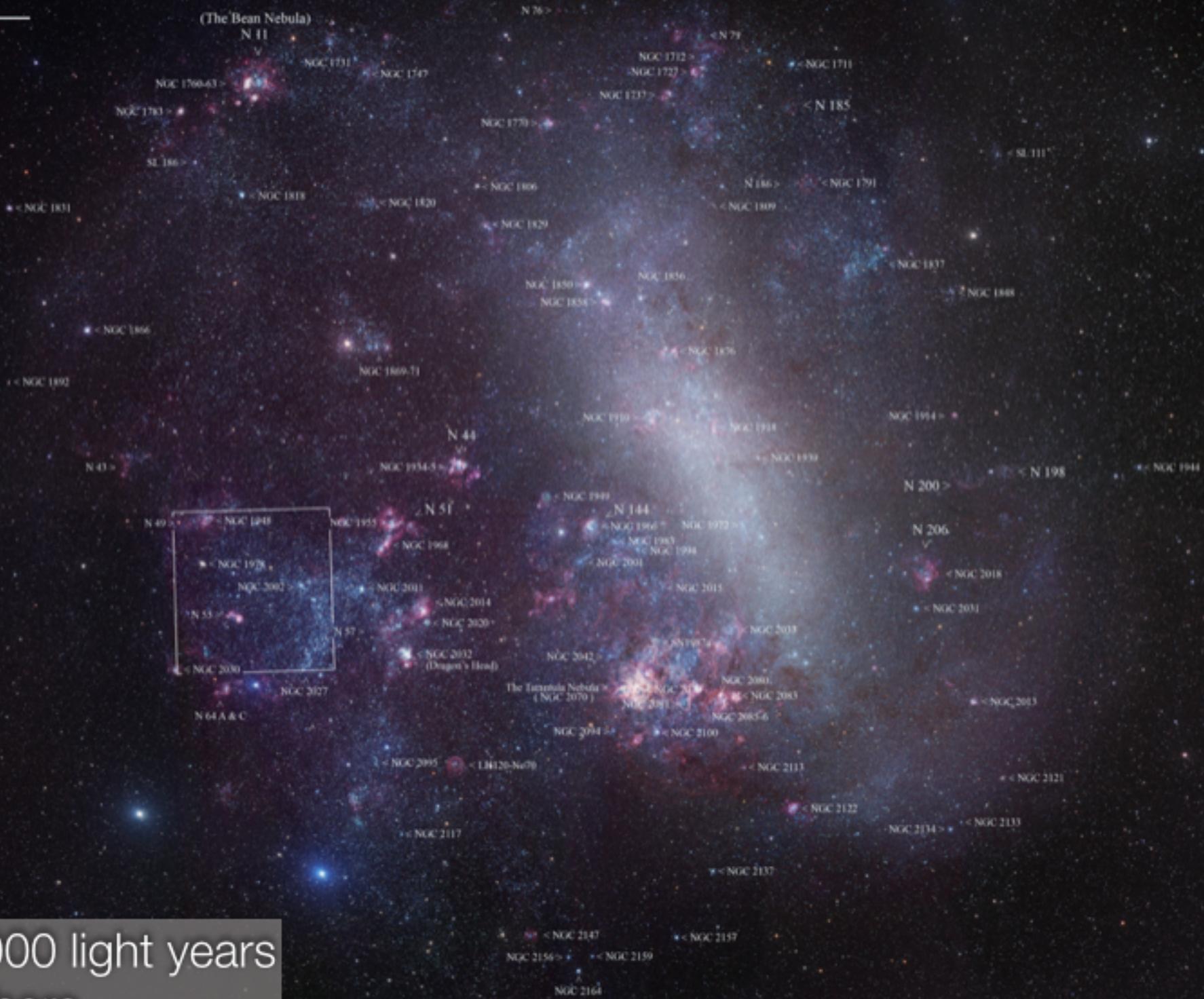




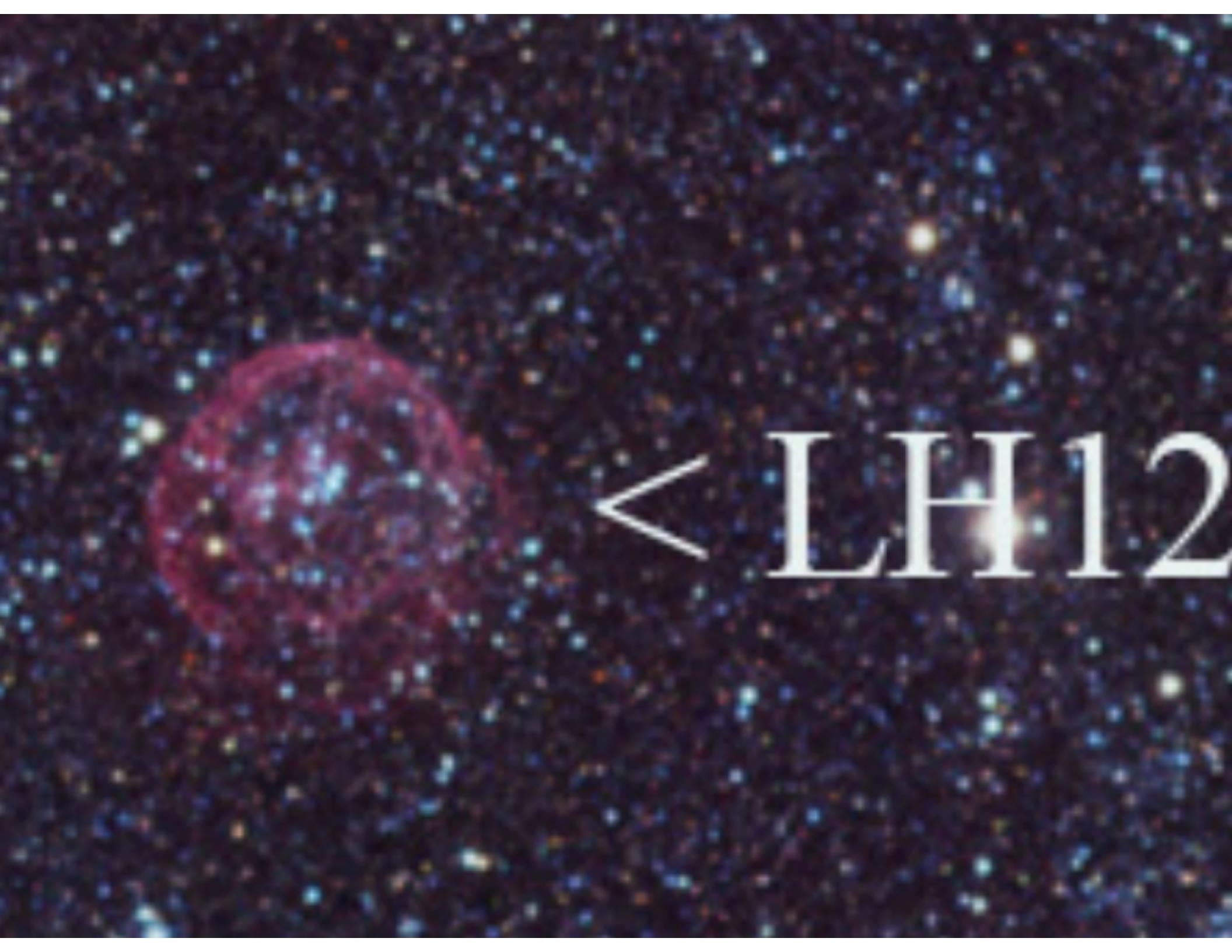
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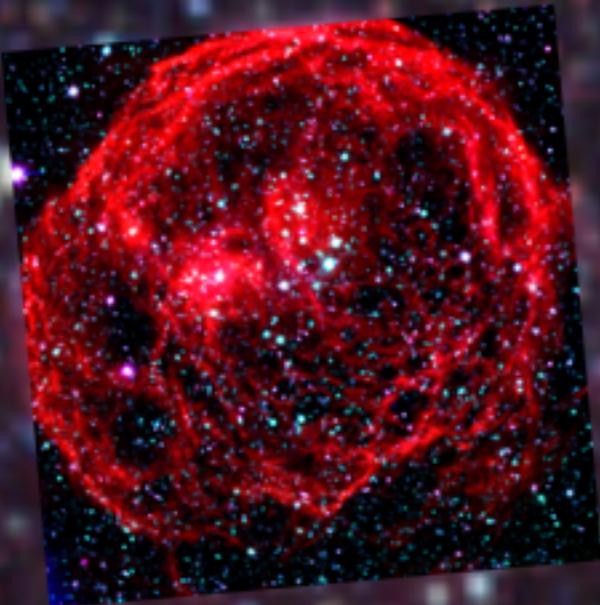
N ←



160,000 light years  
from here



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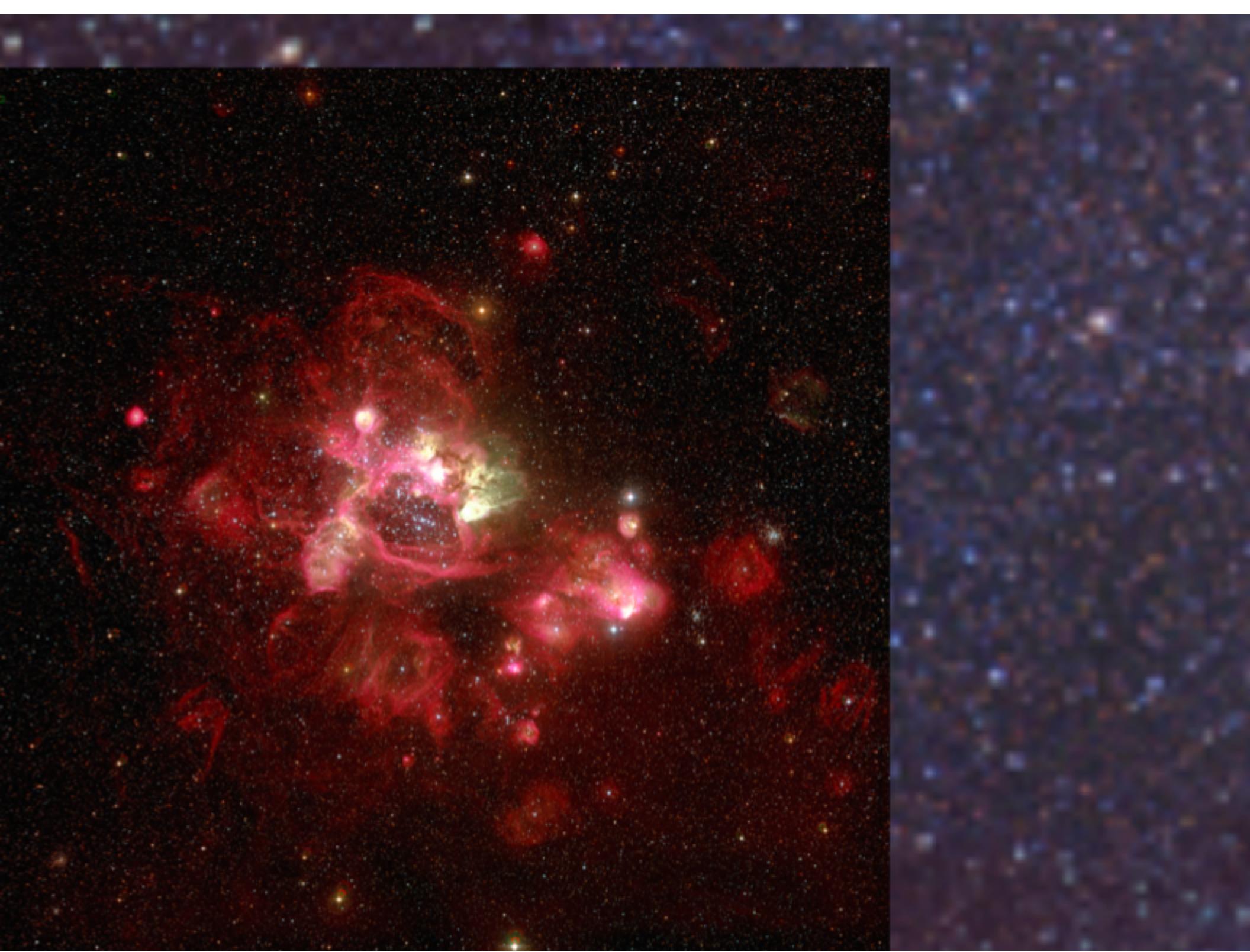
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4-5

160,000 light years  
from here

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(The Bean Nebula)

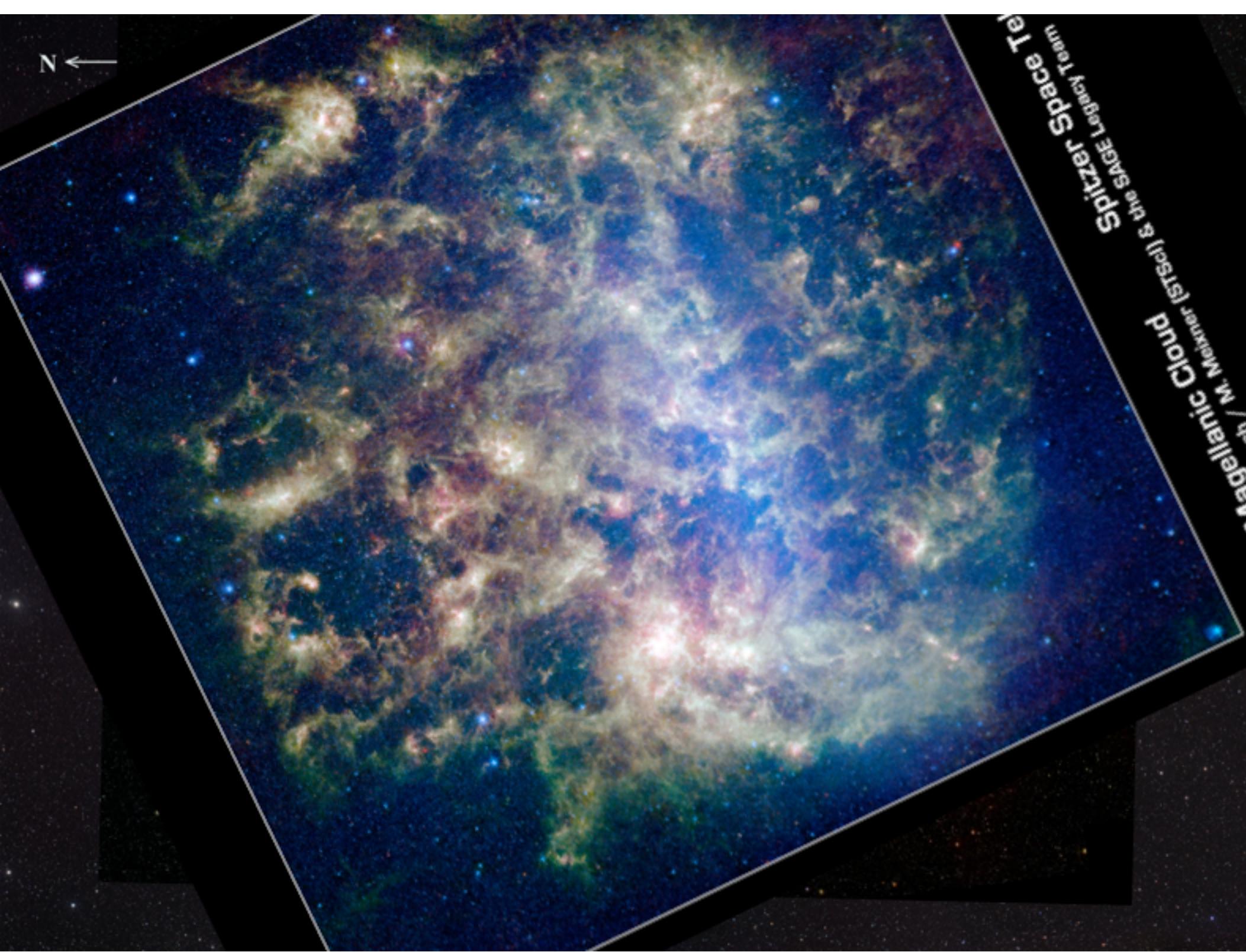
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N ←



NGC 2164



Spitzer Space Team  
SAGE Legacy  
Cloud & the  
Magellanic Cloud  
Grebel / M. Meixner

N ←



Spitzer Space Team  
SAGE Legacy  
Cloud  
Magellanic Cloud  
Grebel / M. Meixner

# DM in Galaxy Clusters

Um, wie beobachtet, einen mittleren Dopplereffekt von 1000 km/sec oder mehr zu erhalten, müsste also die mittlere Dichte im Comasystem mindestens 400 mal grösser sein als die auf Grund von Beobachtungen an leuchtender Materie abgeleitete<sup>1)</sup>. Falls sich dies bewahrheiten sollte, würde sich also das überraschende Resultat ergeben, dass dunkle Materie in sehr viel grösserer Dichte vorhanden ist als leuchtende Materie.

Coma Cluster from the Coma Cluster Treasury Survey (HST)  
Distance: ~320 Mlyr,



Fritz Zwicky  
Born: 14-Feb-1898  
Birthplace: Varna, Bulgaria  
Died: 8-Feb-1974  
in: Pasadena, CA  
Cause of death: Heart Failure  
Remains: Buried, Mollis Cemetery, Glarus, Switzerland

Zwicky  
(1933)

In sections iii, iv, and v three new methods for the determination of nebular masses are discussed, each of which makes use of a different fundamental principle of physics. Method iii is based on the *siriel theorem* of classical mechanics. The application of this theorem to the Coma cluster leads to a minimum value  $\bar{M} = 4.5 \times 10^{10} M_{\odot}$  for the average mass of its member nebulae. Method iv calls for the observation among nebulae of certain *gravitational lens* effects.



Zwicky  
(1937)

# Virgo Cluster

Distance: 53 Mlyr

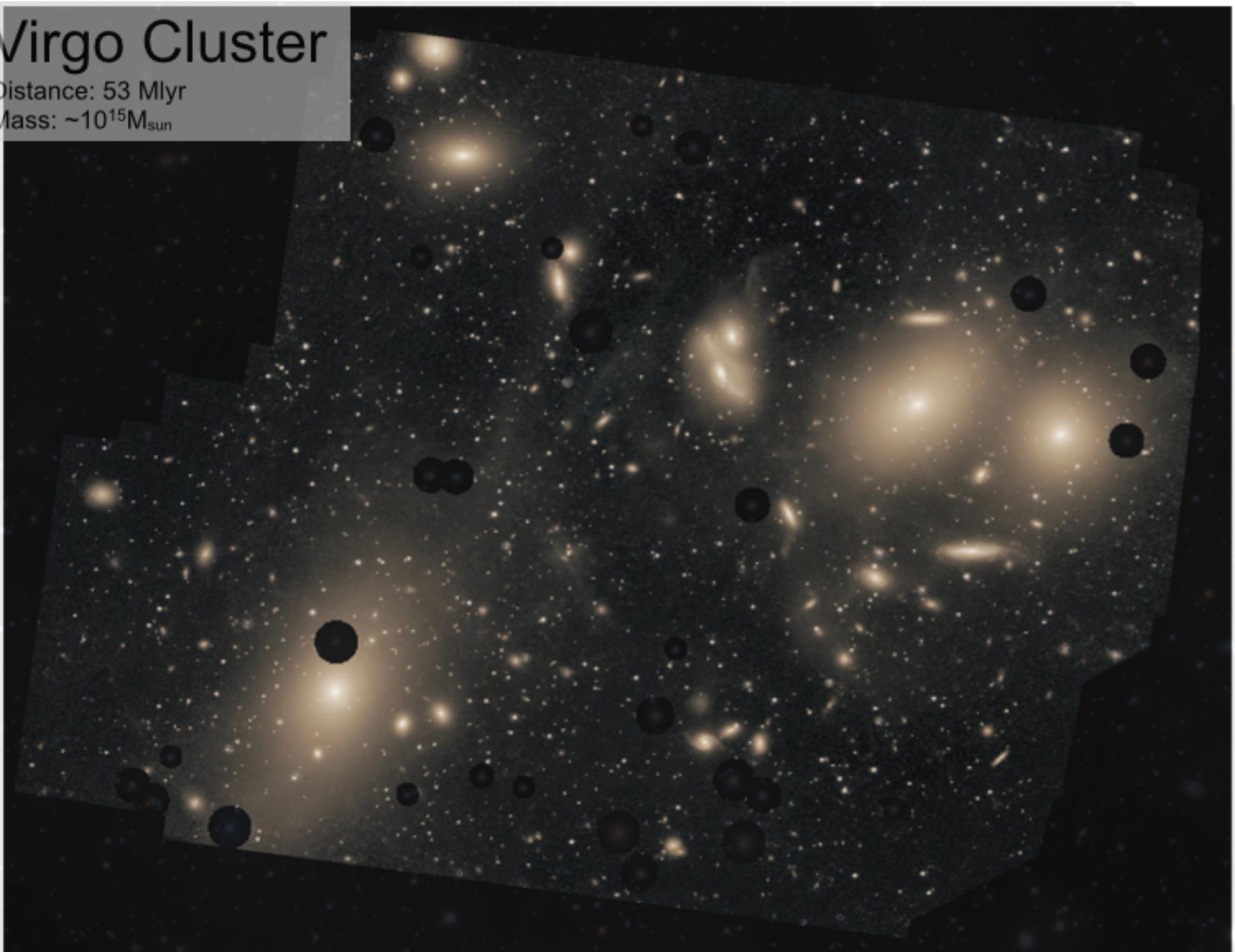
Mass:  $\sim 10^{15} M_{\text{sun}}$



# Virgo Cluster

Distance: 53 Mlyr

Mass:  $\sim 10^{15} M_{\text{sun}}$



## Checklist:

- ✓ Dark Matter exists on Myr scale
- ✓ Distributed differently from stars
- ✓ Also exists in galaxies
- ✓ Many accurate measures of DM

1986BAAS...18R1014L

68.01

### Giant Luminous Arcs in Galaxy Clusters

R. Lynds (KPNO/NOAO), V. Petrosian (Stanford U.)

We announce the existence of a hitherto unknown type of spatially coherent extragalactic structure having, in the two most compelling known examples, the common properties: location in clusters of galaxies, narrow arc-like shape, enormous length, and situation of center of curvature toward both a cD galaxy and the apparent center of gravity of the cluster. The arcs are in excess of 100 Kpc in length, have luminosities roughly comparable with those of giant E galaxies, and are distinctly bluer than E galaxies - especially so in one case. Interpretations of the nature of the arcs are discussed within the framework of available data.

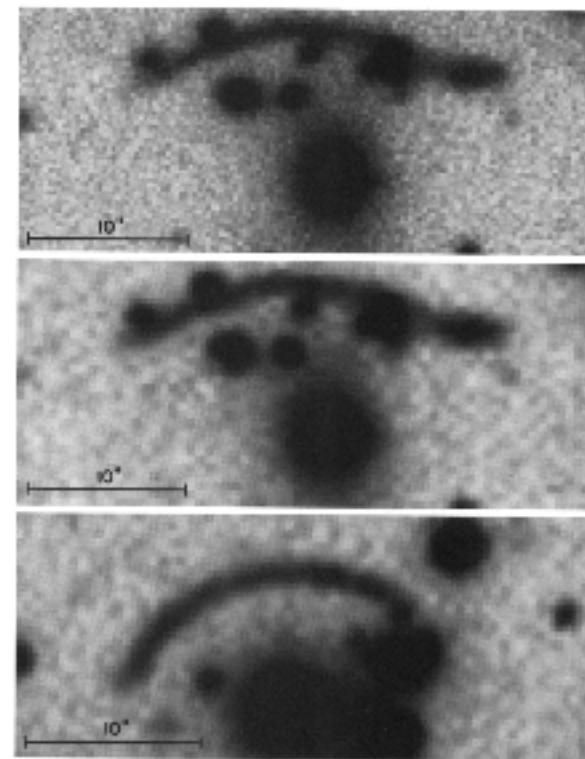


Fig. 1. Three panels of the luminous arcs in the Abell 1656 cluster. The top panel shows a broad, diffuse arc; the middle panel shows a narrower, more concentrated arc; the bottom panel shows another view of the same or a similar arc. Scale bars are indicated in each panel.

# Gravitational Lensing

## Checklist:

- ✓ Dark Matter exists on Mlyr scale
- ✓ Distributed differently from stars
- ✓ Also exists in galaxies
- ✓ Many accurate measures of DM

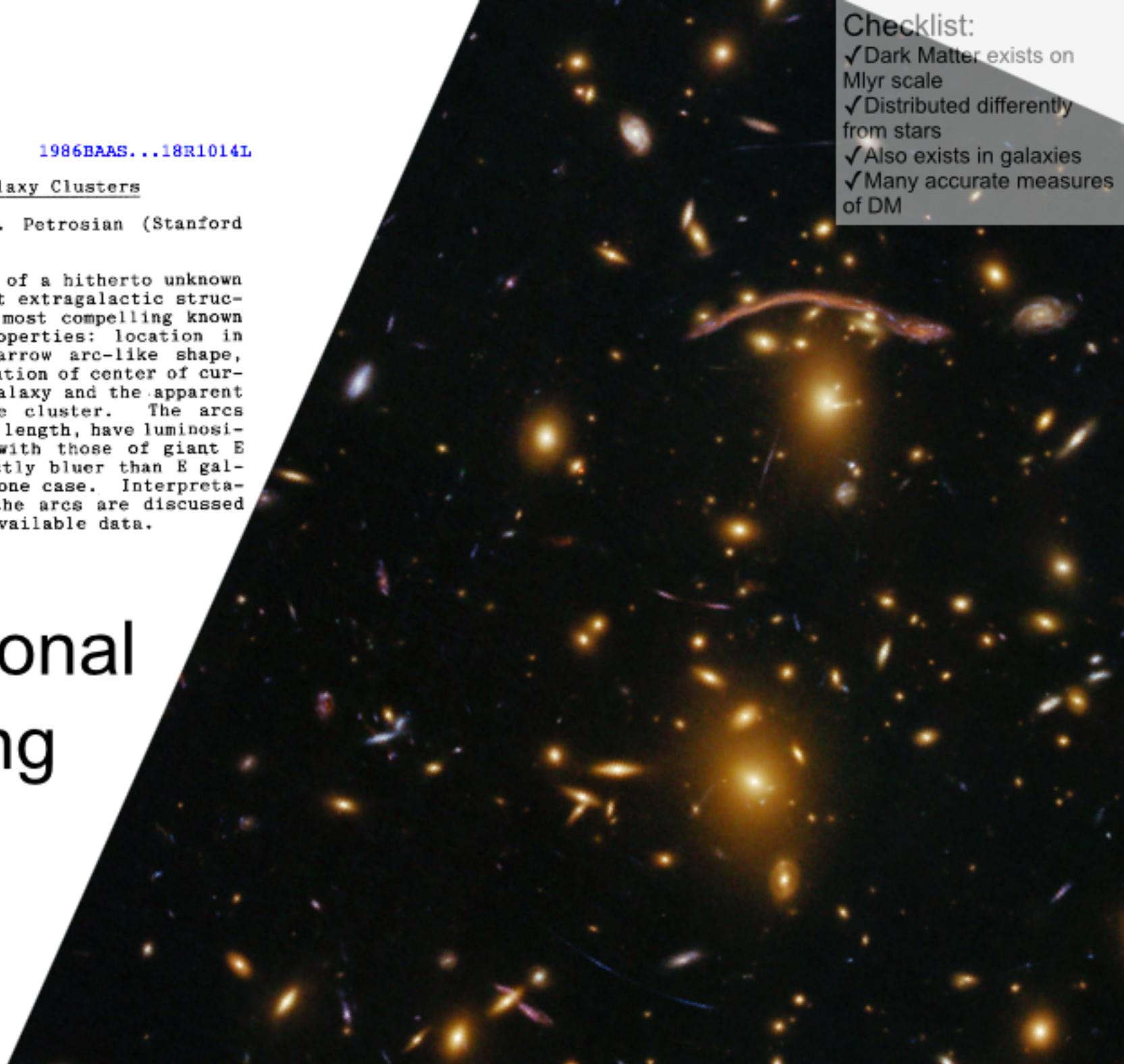
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# Gravitational Lensing

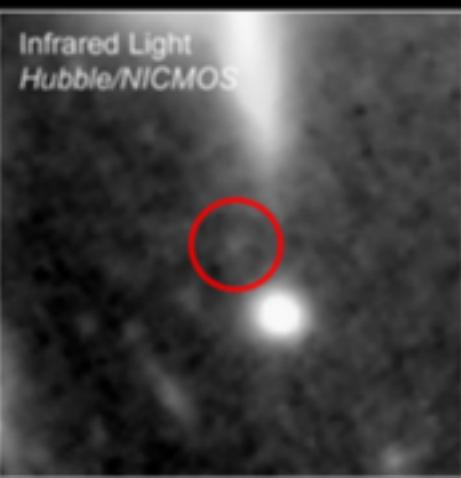
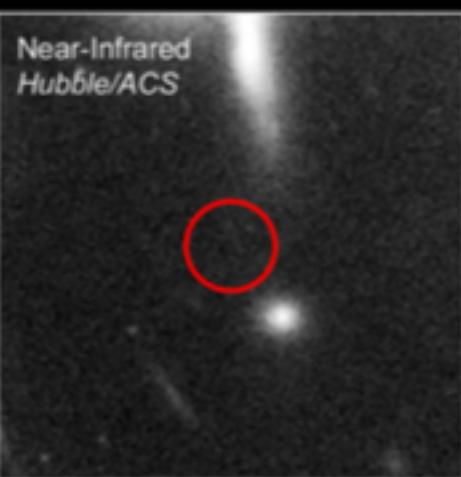
# Dark Matter “seen” with your naked eye

Abell 2218

Distance:

2,3Glyr

Redshift: 0.18



Strong  
Gravitational  
Lensing

### Checklist:

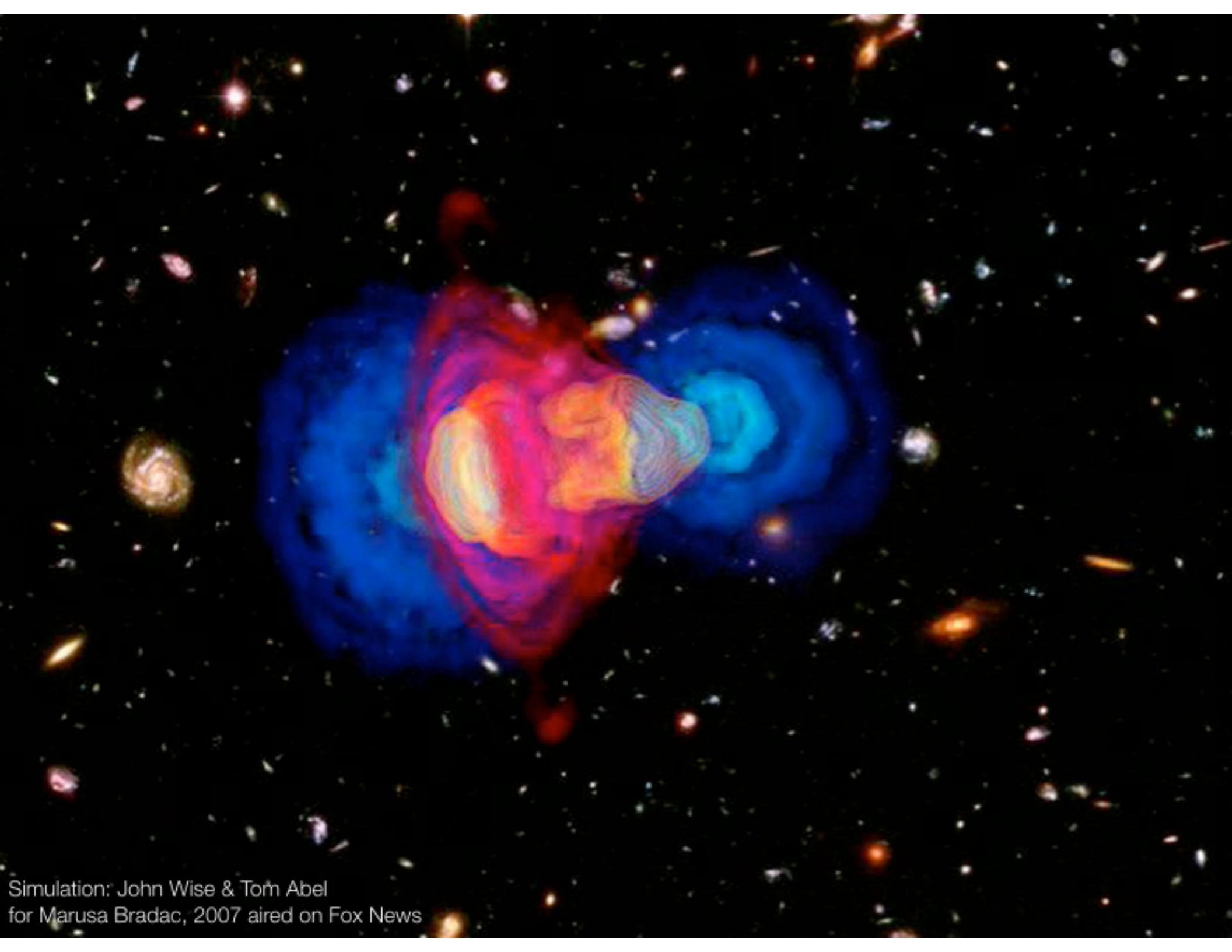
- ✓ Dark Matter exists on Myr scale
- ✓ Distributed differently from stars
- ✓ Also exists in galaxies
- ✓ Many accurate measures of DM
- ✓ DM interacts weakly with gas and itself.  
Collisionless?

## Dark matter & Gravity dominate on large scales

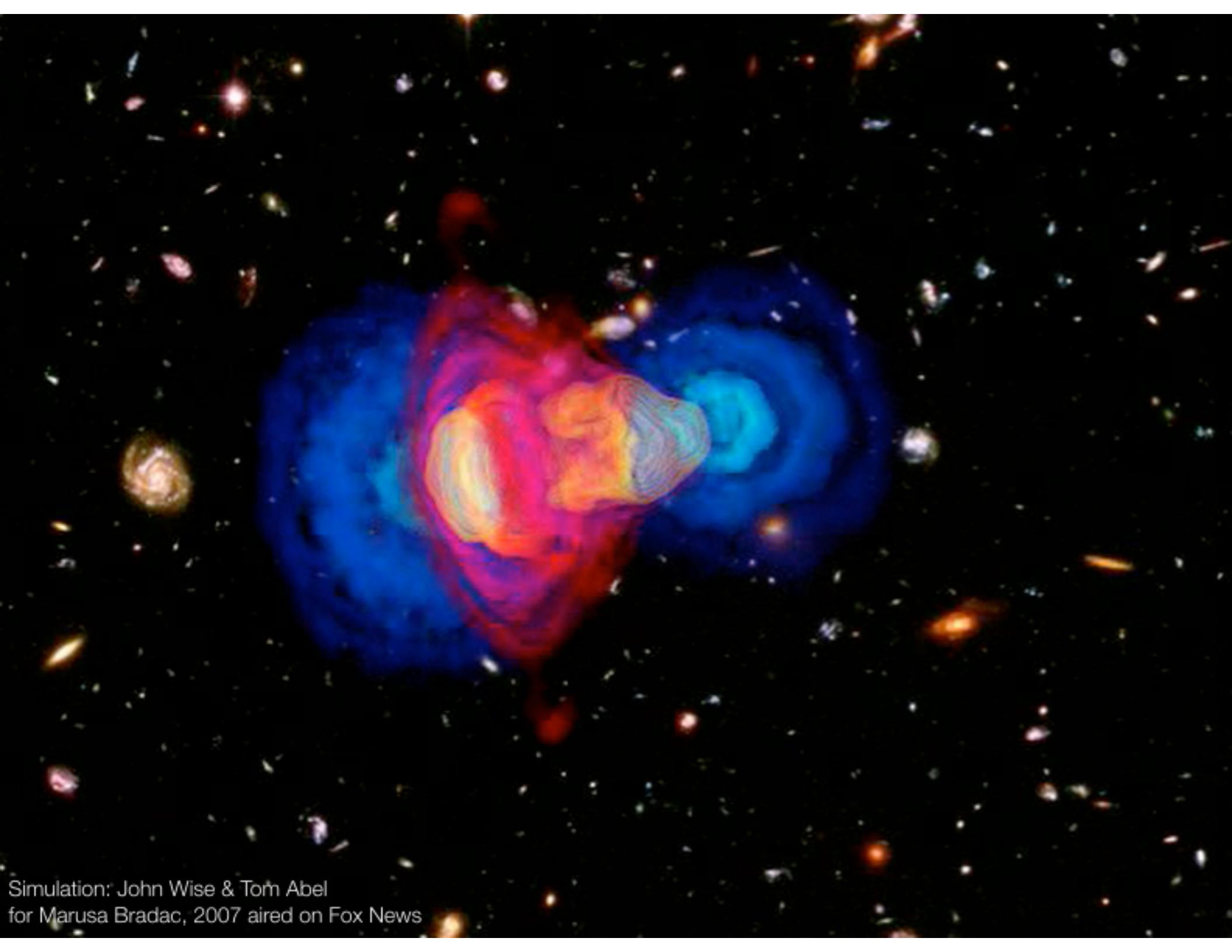


Bullet Cluster  
collision speed 2600 km/s  
mass ratio ~1/10  
Cluster: 1E0657-56  
~ 10 million lyrs across  
~ 10 billion lyrs from earth

Clowe, Bradac, et al. 2007



Simulation: John Wise & Tom Abel  
for Marusa Bradac, 2007 aired on Fox News

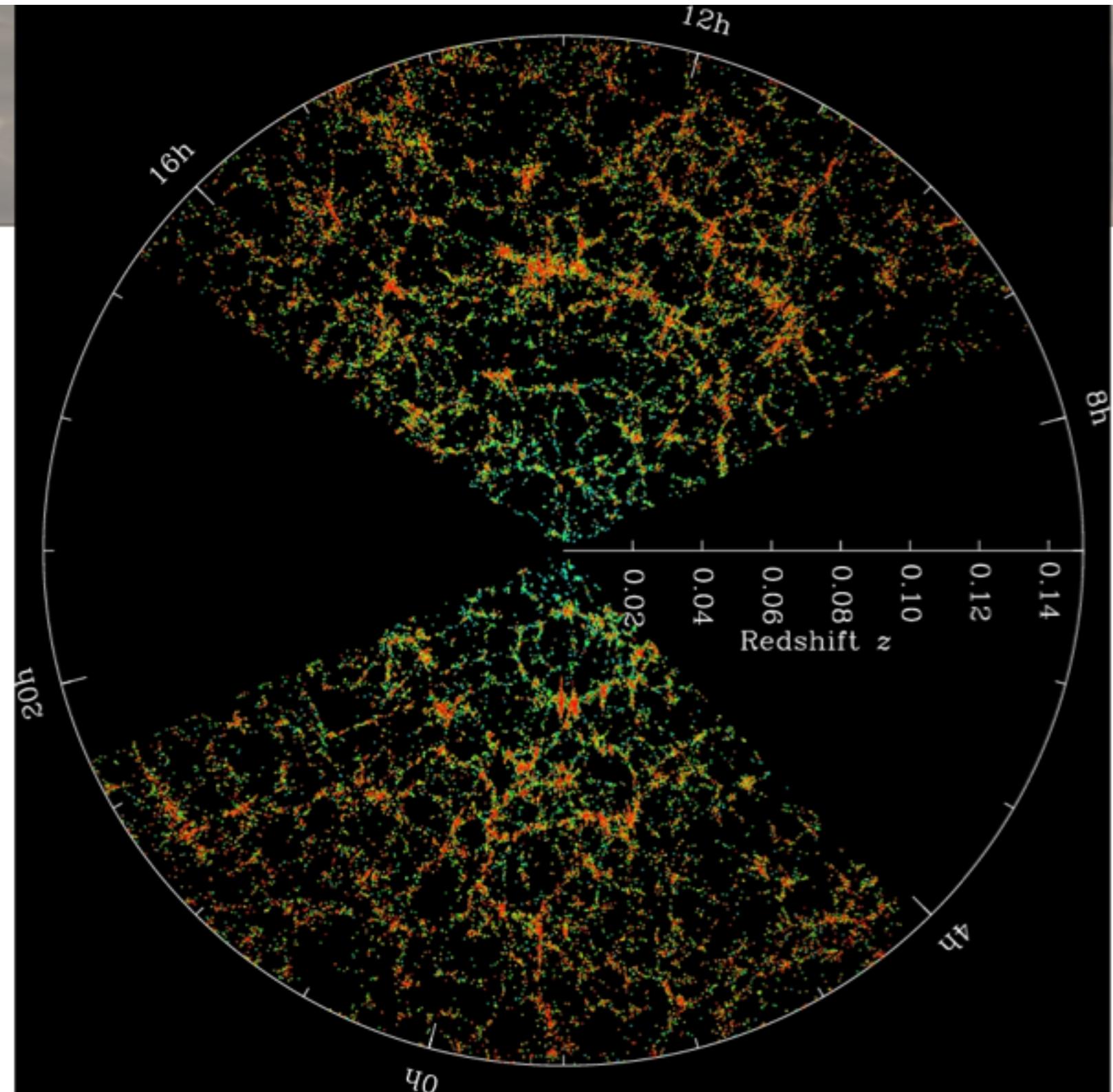


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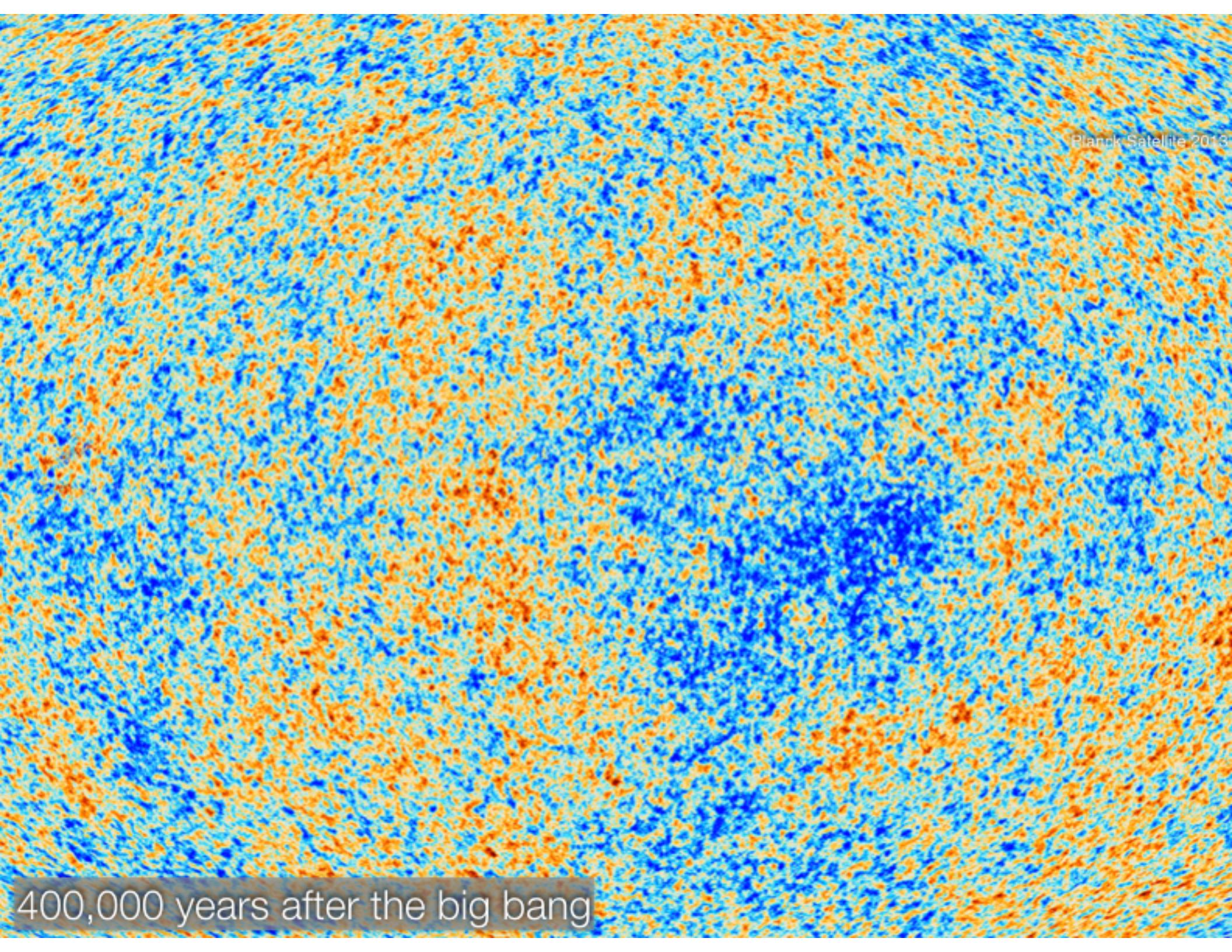
# SDSS

Slices through the SDSS 3-dimensional map of the distribution of galaxies. Earth is at the center, and each point represents a galaxy, typically containing about 100 billion stars. Galaxies are colored according to the ages of their stars, with the redder, more strongly clustered points showing galaxies that are made of older stars. The outer circle is at a distance of two billion light years. The region between the wedges was not mapped by the SDSS because dust in our own Galaxy obscures the view of the distant universe in these directions. Both slices contain all galaxies within -1.25 and 1.25 degrees declination.

Credit: M. Blanton and the Sloan Digital Sky Survey.



13.8 billion years ago

A detailed map of the Cosmic Microwave Background (CMB) radiation, showing temperature fluctuations across the sky. The fluctuations are represented by a color gradient from blue (low temperature) to red (high temperature), with most of the sky appearing in shades of blue and orange. The Planck map shows a complex pattern of small-scale fluctuations and larger, more prominent features like the Great Attractor.

Planck Satellite 2013

400,000 years after the big bang

Viz: Ralf Kaehler & Tom Abel  
For "Dark Universe" planetarium show

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Showing continuously in the Big Bang Theater of the American Natural History Museum in New York

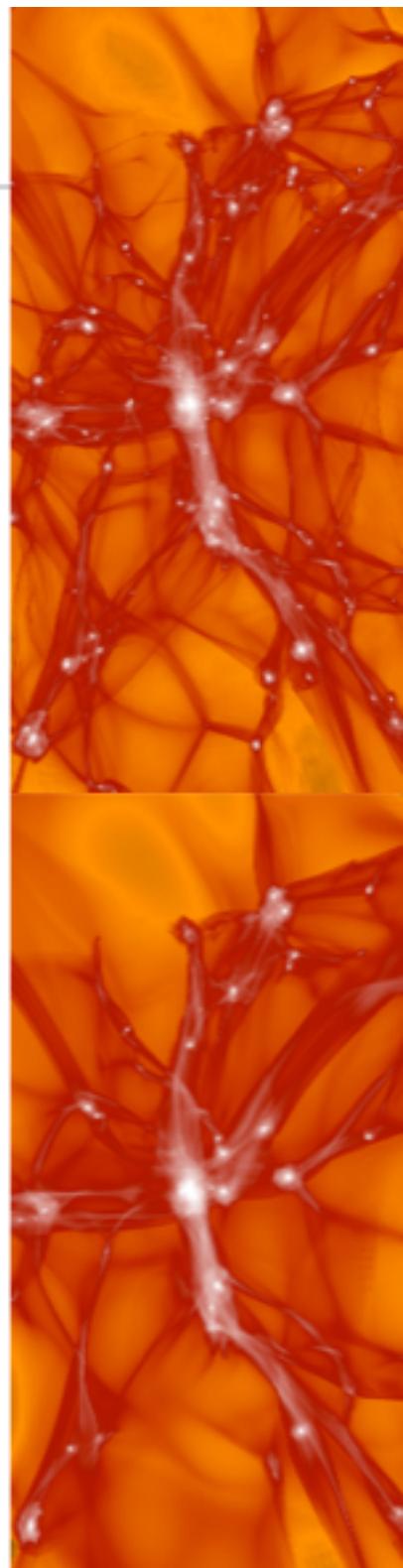
Viz: Ralf Kaehler & Tom Abel

Showing continuously in the Big Bang Theater of the American Natural History Museum in New York

# Cosmological N-body simulations

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- Used to make predictions about the distribution of dark matter in the Universe
- Key results
  - Galaxies are arranged in cosmic web of voids/sheets/filaments/halos
  - Universal spherical Dark Matter density profile (NFW) [not understood from analytical arguments]
  - Predicted mass functions of halos and their clustering and velocity statistics
- Primary tool to study observational consequences of LCDM
  - initial conditions: warm vs cold DM, Gaussian vs non-Gaussian
  - sensitivity on global cosmological parameters such as the total matter content and amount of dark energy, etc.
  - Gravitational Lensing signatures

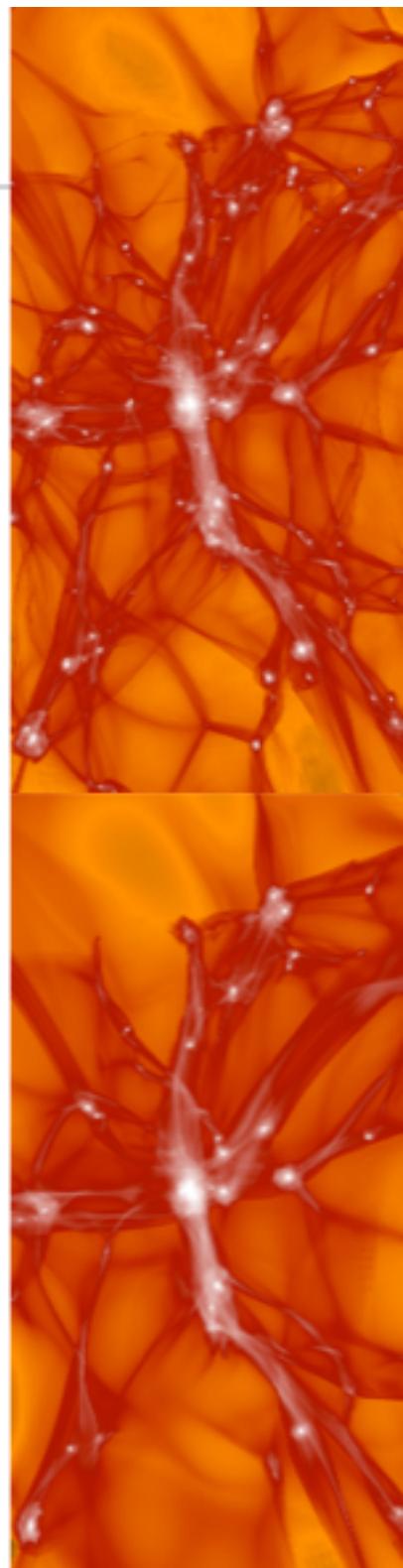
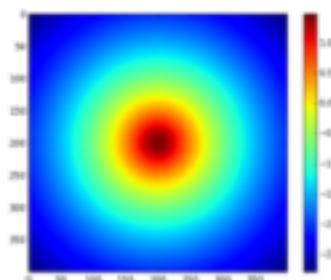


# Cosmological N-body simulations

$$\dot{\mathbf{x}} = \mathbf{v}(t) \quad \dot{\mathbf{v}}_i = - \sum_{i \neq j}^N G m_i m_j \frac{(\mathbf{x}_j - \mathbf{x}_i)}{|\mathbf{x}_j - \mathbf{x}_i|^3}$$

- All modern cosmological simulation codes only differ in how they accelerate the computation of the sum over all particles to obtain the net force
- End result are simply the positions and velocities of all particles
- Limit N goes to infinity must give correct answer, right?
- Softening of forces (add  $\epsilon^2$  in denominator) avoids singularities.

$$\dot{\mathbf{v}}_i = - \sum_{i \neq j}^N G m_i m_j \frac{(\mathbf{x}_j - \mathbf{x}_i)}{(|\mathbf{x}_j - \mathbf{x}_i|^2 + \epsilon^2)^{3/2}}$$



# Some public Cosmological N-body codes

- Enzo: Adaptive mesh DM+Hydro+MHD+Radiation etc.  
<http://enzo-project.org>
- Gadget: Popular DM only code and SPH hydro:  
<http://www.mpa-garching.mpg.de/gadget/>
- Ramses: Adaptive mesh cosmology code  
<http://www.ics.uzh.ch/~teyssier/ramses/RAMSES.html>
- Initial conditions generator:  
<http://www.phys.ethz.ch/~hahn/MUSIC/>
- Yt: Simulation output analysis  
<http://yt-project.org>



# Summary

- We have now an appreciation of
  - how gravitational instability forms structure in the Universe
  - rough idea how cosmological N-body codes work
  - some of the challenges in studying the fully non-linear dynamics of a self-gravitating collisionless fluid