



# Cosmic Rays

**Advanced FLUKA Course**

# Galactic Cosmic Rays

## Composition:

90% protons, 9% Helium,  
< 1% Ions (*particles*)  
64% protons, 25% Helium,  
11% Ions (*nucleons*)

## Spectrum:

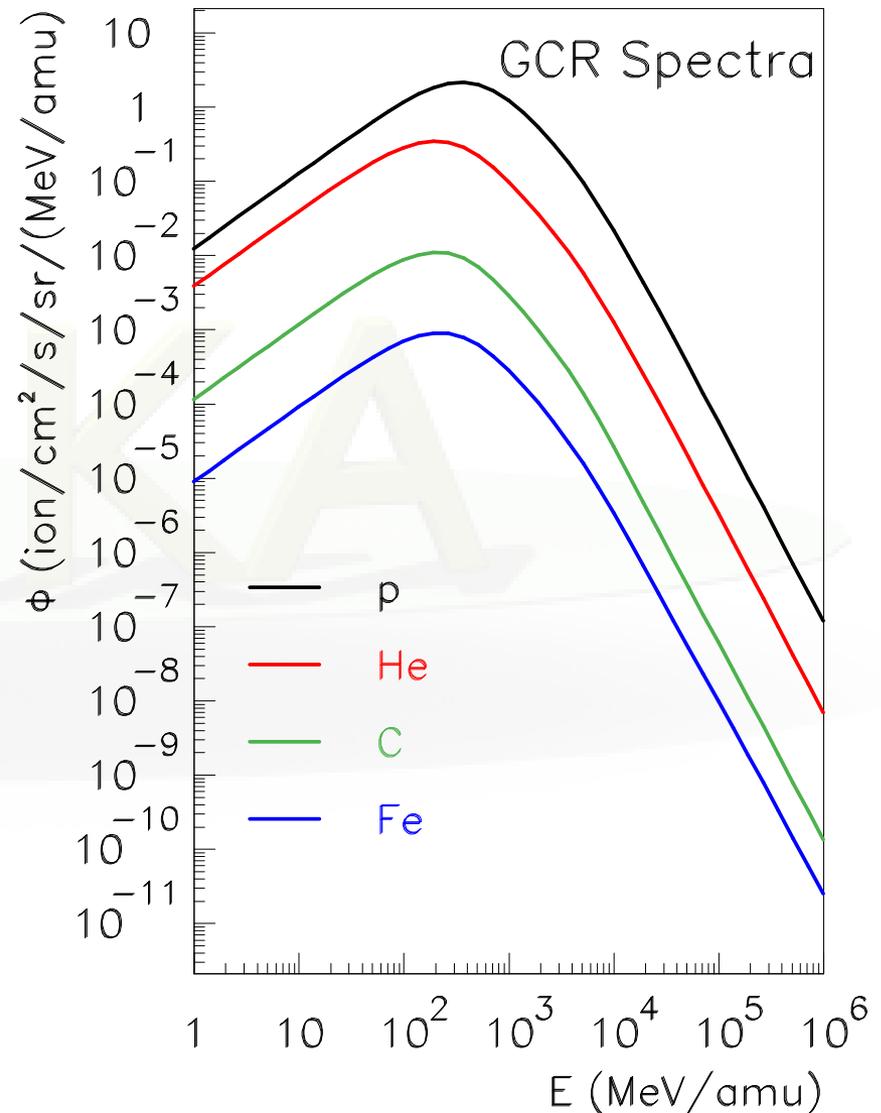
broad spectrum, peaks around  
1 GeV/n

## Intensity:

( $E > 10$  MeV/n)  $\sim 5$  p/(cm<sup>2</sup> s)  
@ Solar Min.

## Dose/Dose Equivalent:

$\sim 0.4$  mGy/d, 1 mSv/d (no  
geomagnetic cut off)



# GCR: How To Calculate?

- ❑ the determination of the **spectrum** and **composition** of cosmic rays at the **local interstellar medium**
- ❑ the determination of changing conditions in the **solar wind magnetic field** and the resulting **interaction** with the **inward flow of galactic cosmic rays** from the local interstellar medium
- ❑ the determination of the **trajectories** of cosmic rays through the Earth's **geomagnetic field**
- ❑ a realistic description of the **Earth's atmosphere**
- ❑ the **transport of the surviving incident cosmic rays** through the **Earth's atmosphere** to various depths.

# GCR in FLUKA

- Primary **cosmic ray spectra** and interplanetary modulation according to measured **solar activity** on a day-to-day basis for past years are available
- **Geomagnetic effects** are implemented with full multipole expansion
- **Extensive benchmarks** against available muon and hadron measurement data
- Used by **several laboratories** (CEA Saclay, Frascati, Siegen, Bartol, Houston, GSF,...) for simulating the radiation fields generated by **cosmic rays in the atmosphere and/or inside spacecrafts and satellites**

# GCR: How To Calculate with FLUKA

The Galactic Cosmic Ray (**GCR**) component of the cosmic ray flux can be simulated up to 30 TeV/nucleon (or **1000 TeV/n** when DPMJET is linked)

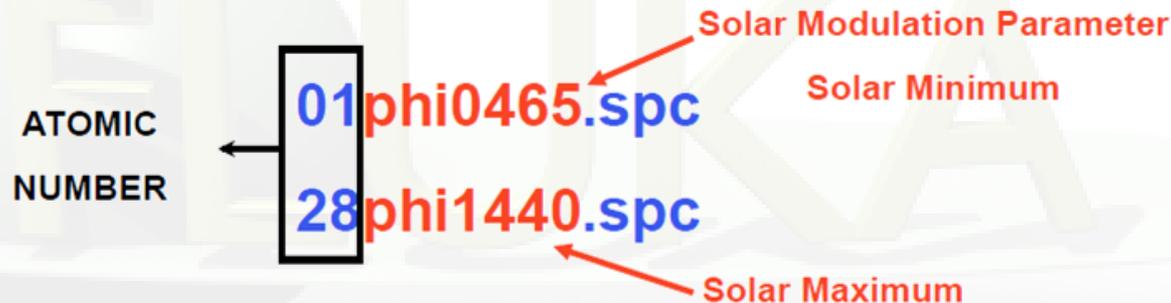
The following **general options** are available concerning the simulation of cosmic ray interactions in FLUKA, wrt ion interactions:

- **Superposition model:** in this approach (**All-Nucleon Spectrum**) primary **nuclei are split into equivalent independent nucleons**. See card **PHYSICS** with **SDUM = IONSPLITti**
- **DPMJET (rQMD) interaction models** (**All-Particle Spectrum**): these models **simulate nucleus-nucleus collisions (DPMJET above 5 GeV/nucleon, rQMD below)**. In case DPMJET is chosen for cosmic ray application, for very high energies ( $\sim 10^{16}$ eV) it is recommended to use the **DPMJET-II.5** version instead DPMJET-III (linking with the script `$FLUPRO/flutil/ldpm2qmd`)

*The DPMJET and the superposition model can also be used together, by setting the respective energy ranges with the **PHYSICS** card.*

# The All-Particle spectrum

- ❑ The ion composition of the galactic flux is derived from a code (Badhwar & O'Neil 1996) which considers all elemental groups from  $Z = 1$  to  $Z = 28$  (Ni). The spectrum is modified to follow recent data sets from the AMS and BESS experiments up to 100 GeV, according to the so-called **ICRC2001** fit.
- ❑ The spectrum components are written in 28 files:



- ❑ These spectra are without geomagnetic cutoff. They are used together with a few analytical calculations of the rigidity cutoff, according to different descriptions of the Earth geomagnetic field (see later)

**Spectra of Solar Particle Events of Jan 20<sup>th</sup>, 2005 and Oct 28<sup>th</sup> 2003 are also available**

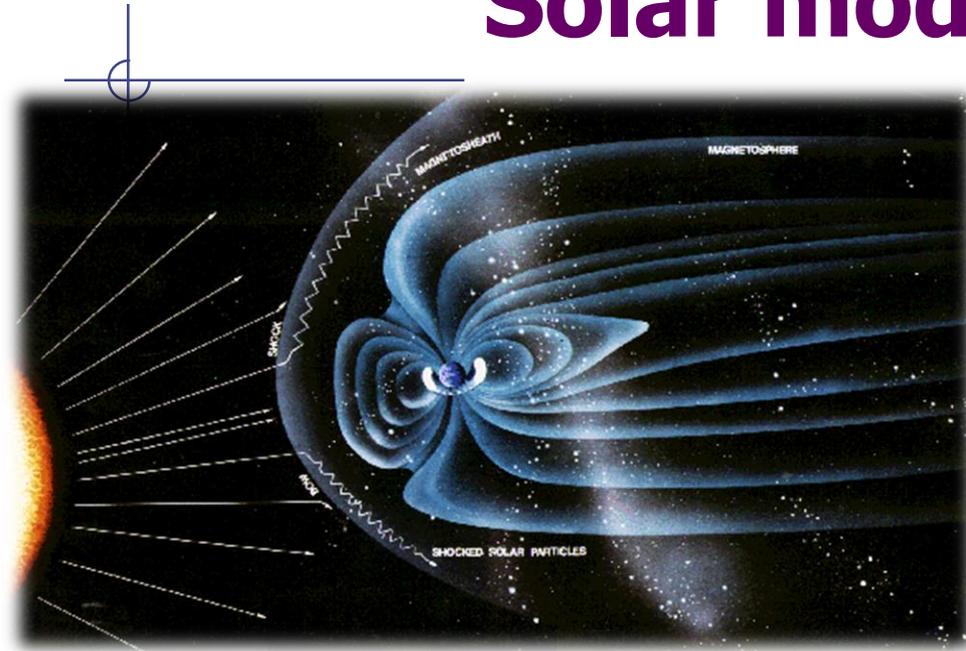
# The ALL-Nucleon Spectrum (1)

- ❑ Based on a modified fit of the **All-Nucleon flux proposed by the Bartol group**, using the All-Particle Spectrum up to 100 GeV and data published in ICRC 2003
- ❑ For the **proton component at energies larger than 100 GeV**, using the normalization obtained at 100 GeV, a spectral index  $\gamma = -2.71$  is assumed
- ❑ A spectral index  $\gamma = -3.11$  is assumed **above the knee at 3000 TeV**
- ❑ For what concerns the **He component**,  $\gamma = -2.59$  is used **above 100 GeV** and a charge-dependent knee is assumed according to the rule:  $E_{\text{nucleon}} = Z * 3000 \text{ TeV}/A$
- ❑ Higher Z components have been grouped in **CNO, MgSi** and **Fe** sets and treated using an All-Particle spectrum with the above mentioned charge-dependent knee parameterization

# The ALL-Nucleon Spectrum (2)

- ❑ Fluxes are read from a file named "**allnucok.dat**" giving the **total energy (GeV)**, the **fluxes (E.dN/dE)** and the **neutron/proton ratios**
- ❑ This option ("All Nucleon Flux") is chosen with command **SPECSOUR** and **SDUM = GCR-ALLF**
- ❑ The user can decide whether
  - to **sample neutrons and protons** from the file and to transport them **using the superposition model**
  - or
  - to consider all **neutrons as being bound in alpha particles** and to **transport protons and alphas (better for magnetic field effects)**

# Solar modulation (1)



- The deviation from the power law, observed below 10 GeV, is a consequence of the influence of the **solar wind** called **solar modulation**.
- Flux intensity in this energy range is **anti-correlated to the solar activity** and follows the **sun-spot 11-year cycle**
- The correlation between the solar activity and the modulation of the cosmic rays flux has been studied by monitoring the flux of atmospheric neutrons.

- ❑ A flux of low energy neutrons is produced in the interaction of primary CRs with the atmosphere and it is mostly due to low energy primaries (1-20 GeV), due to the rapid fall of the primary flux intensity with energy → those neutrons can be detected at ground level
- ❑ One assumes that far from the solar system there exists an unmodified flux called **Local Interstellar Spectrum**, which is modified within the solar system by the interaction with the solar wind. This interaction is well described by the **Fokker-Planck diffusion equation**, that can be obtained describing the solar wind by a set of magnetic irregularities, and considering these irregularities as perfect elastic scattering centers
- ❑ For energies above 100 MeV this equation can be solved using the "**Force Field Approximation**"

# Solar modulation (2)

- According to the **Force Field Approximation**, at a given distance from the Sun, for example at 1 a.u., the population of CRs at energy  $E_{interstellar}$  is shifted at the (lower) energy  $E_0$  as in an energy loss mechanism due to a potential  $V$ :

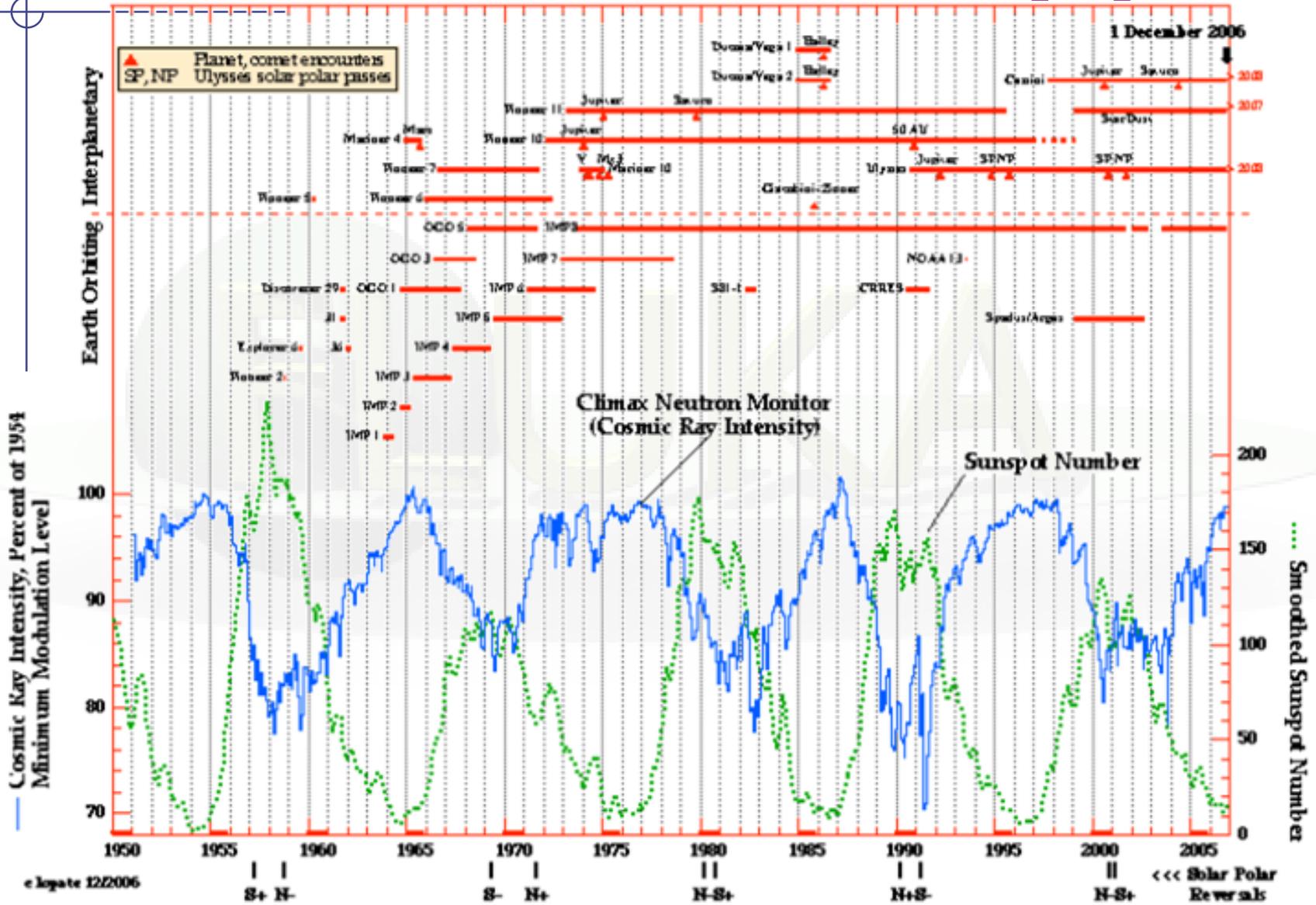
$$E_0 = E_{interstellar} + Z \times V_{solarwind}(t)$$

- The solar wind potential at a given distance from the Sun depends on only one parameter, the time:  $V = V(t)$ . So it doesn't matter what the interstellar flux is: given a flux on the Earth at a time  $t$ , one can find the flux at another time just from the relative variation of the solar wind potential  $V$ .
- In FLUKA, an offline code uses an algorithm which takes into account
  - ❑ either a given  $V$  value expressing the effect of the interplanetary modulation of the local interstellar spectrum
  - ❑ or the counting rate of the **CLIMAX** neutron monitor to provide the flux prediction at a specific date if available

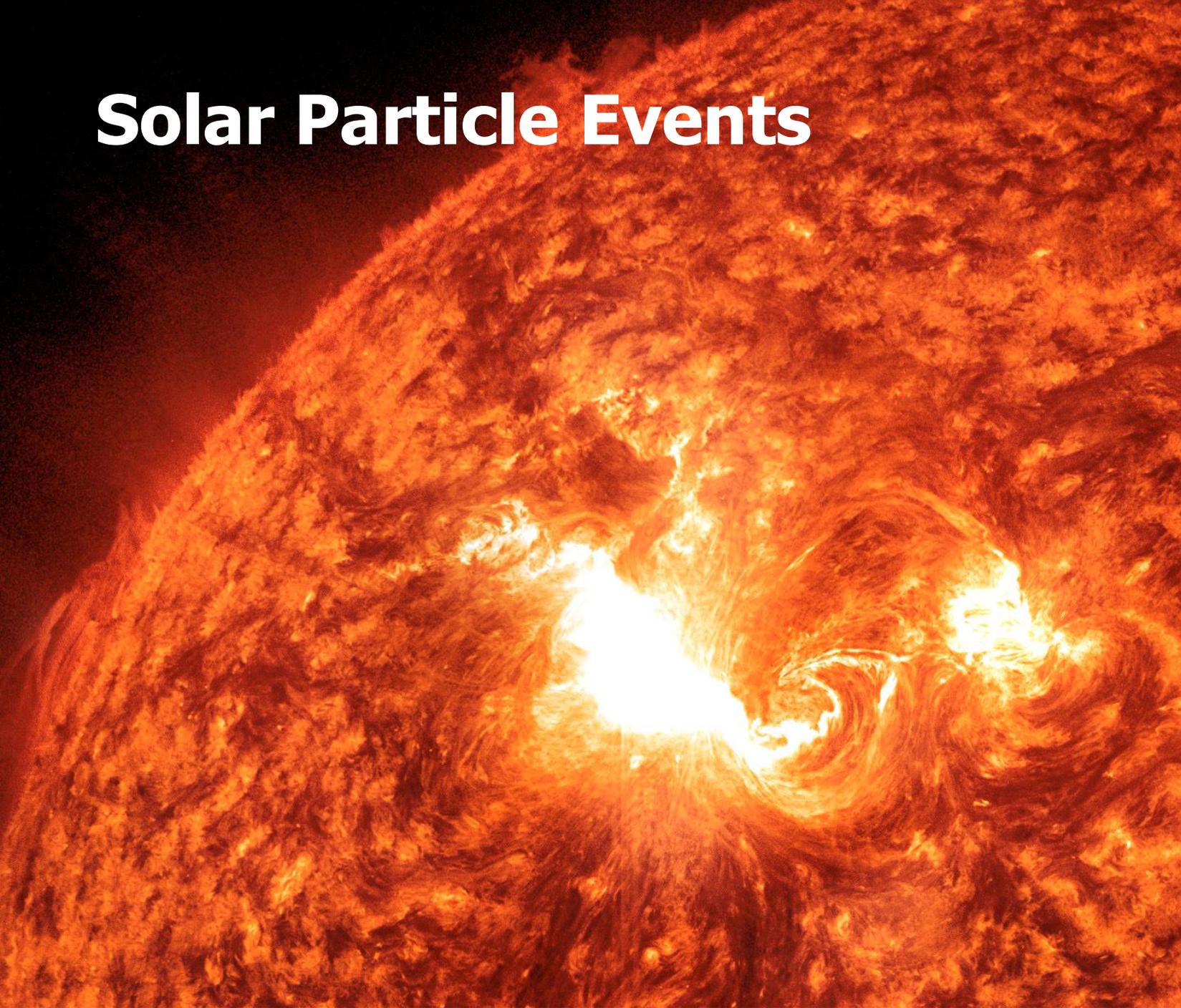
***The model is not a description of the processes and of the way in which they occur, but reasonably predicts the GCR modulation at Earth.***



# Solar modulation (3)



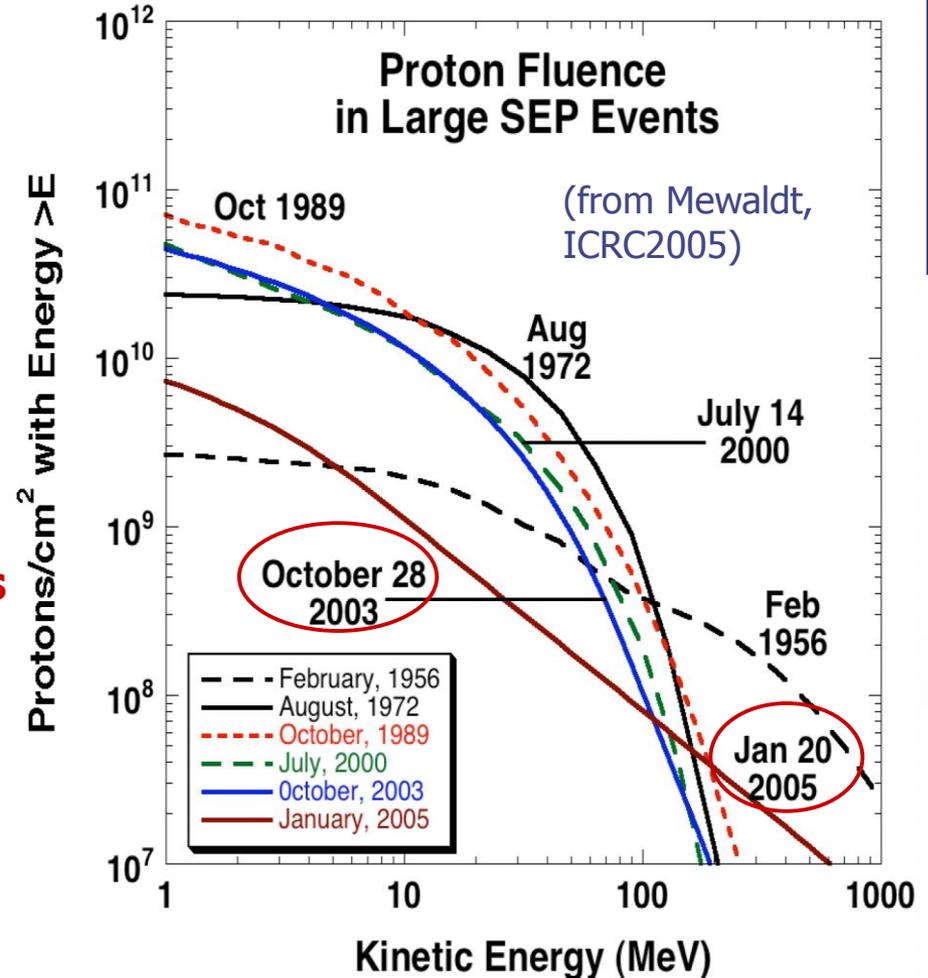
# Solar Particle Events



**A  
Solar  
Flare**

# Solar Energetic Particle (SEP) events

- Integrated fluence: up to  **$10^{11}$**  (**nucleons/cm<sup>2</sup>**),  $E > 1$  MeV / n
- Large **variations** in **spectra**
- Variable composition: mostly **protons (~90%)** and  **$\alpha$ 's (~9%)**, but ions up to Iron are not negligible
- Variable **duration**, from **hours to days**
- **Rise** time from **minutes to hours**
- **Dose equivalent** up to  **$\sim$ Sv**, highly dependent on organ, shielding, and SEP intensity/spectrum
- **Unpredictable** 

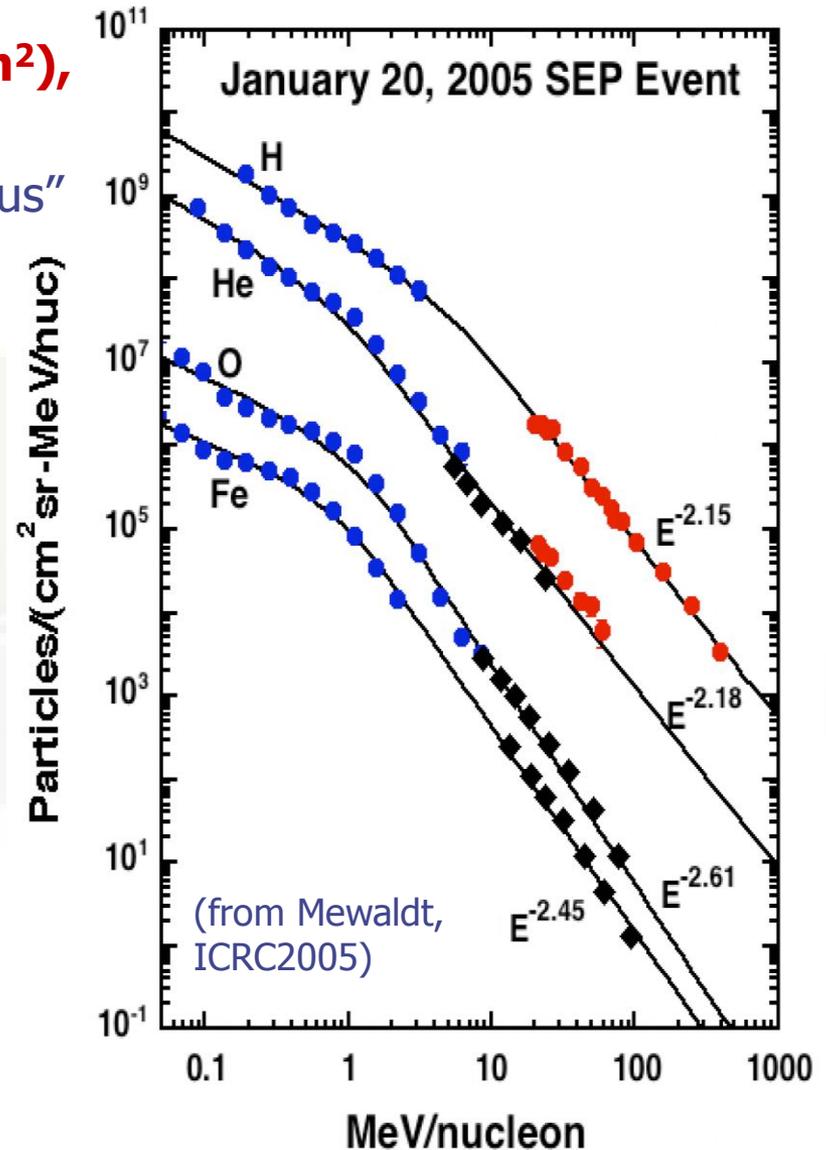


***Nightmare scenarios for (manned) missions beyond Earth low orbits***

# SEP: the “hard” event of 20 Jan 2005

- ❑ Integrated flux:  $\sim 7 \cdot 10^9$  (nucleon/cm<sup>2</sup>),  
E > 1 MeV / n
- ❑ The hardest spectrum after the “famous”  
February 1956 event
- ❑ Detectable increase in ground level  
muons above **5 GeV** !!!
- ❑ Very fast rise time ( $\sim$  minutes)

**Solar Particle Events: lower energies but much higher intensities than Galactic Cosmic Rays**



# Atmospheric model: geometry (1)

## The Earth atmosphere model

- ❑ The FLUKA package makes use of a **density vs. height profile** of atmosphere
- ❑ **An external program** containing a **functional fit to this profile** is used to **generate at the same time** an input geometry file, together with the **data cards for material description** (each atmospheric layer, having its proper density, needs to be assigned a different FLUKA material)
- ❑ The geometry produced, and distributed with the name **atmogeo.cards** is a spherical representation of the **whole Earth atmosphere**.
- ❑ The **material definitions and assignment** contained in the file **atmomat.cards** correspond to the density profile of the U.S. Standard atmosphere. The cards contained in atmomat.cards shall be included by the user in the input file.

# Atmospheric model: geometry (2)

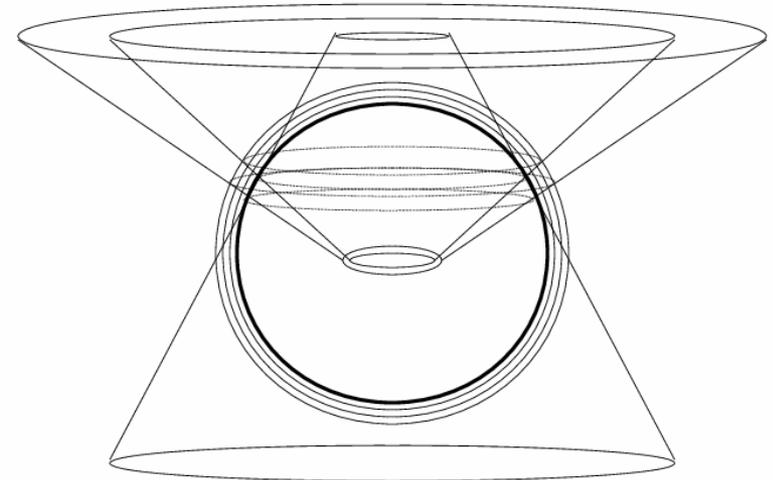
## The Earth atmosphere model

- In addition, the user can specialize this geometry to a given geomagnetic latitude with the help of the **atmloc\_2011.f** auxiliary program. In this way, **the geometry will contain only a slice of the atmosphere**, centered on the given (geomagnetic) latitude
- The local geometry file produced by atmloc\_2011.f is named **atmloc.geo**. The user shall rename this geometry file for further use. More auxiliary files are produced by atmloc\_2011.f:
  - ❑ the file **atmlocmat.cards** contain additional material assignments to be included in the input together with the ones from **atmogeocards**
  - ❑ the file **atmloc.sur** contains data used by FLUKA runtime, and normalization areas

# Atmospheric model: geometry (3)

## Local atmosphere model

- The geometry is built using **two truncated cones** (TRC) whose vertex is in the center of the Earth, the base out of the atmosphere and the height (considering a geographical location in the northern hemisphere) is in the direction of the Earth radius which passes through the North Pole



**The angular span between the two cones contains the atmosphere of interest for the latitude of interest. In addition there is a **third cone** placed in the opposite direction: its vertex is where the other two cones have the base, its base is out of the atmosphere and its height is in the direction of the Earth radius which passes through the South Pole. This cone assures that the **horizon** of the zone of interest is properly included into the calculation**

*A similar geometry can be built for a requested latitude in the southern hemisphere.*

# Atmospheric model: geometry (4)

## Local atmosphere model

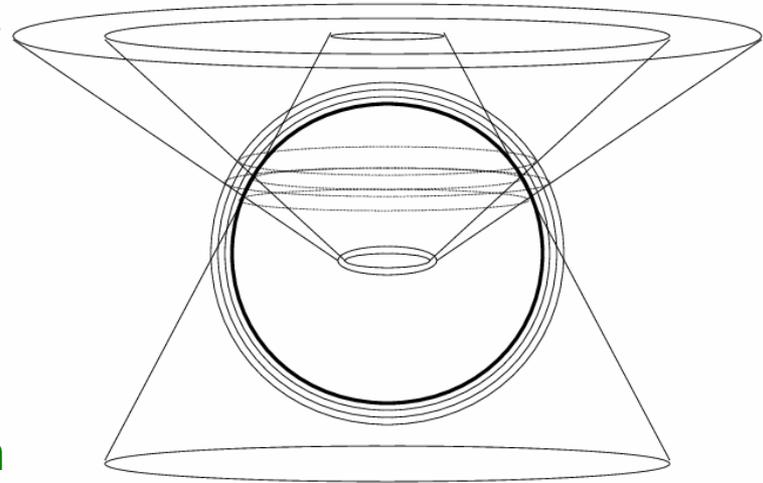
The user builds the complete geometry of the local model, specialized to a given geomagnetic latitude, with the auxiliary program **atmloc\_2011.f**.

The geometry (file **atmloc.geo**), containing only one slice of the atmosphere centered on the given position, will be made of:

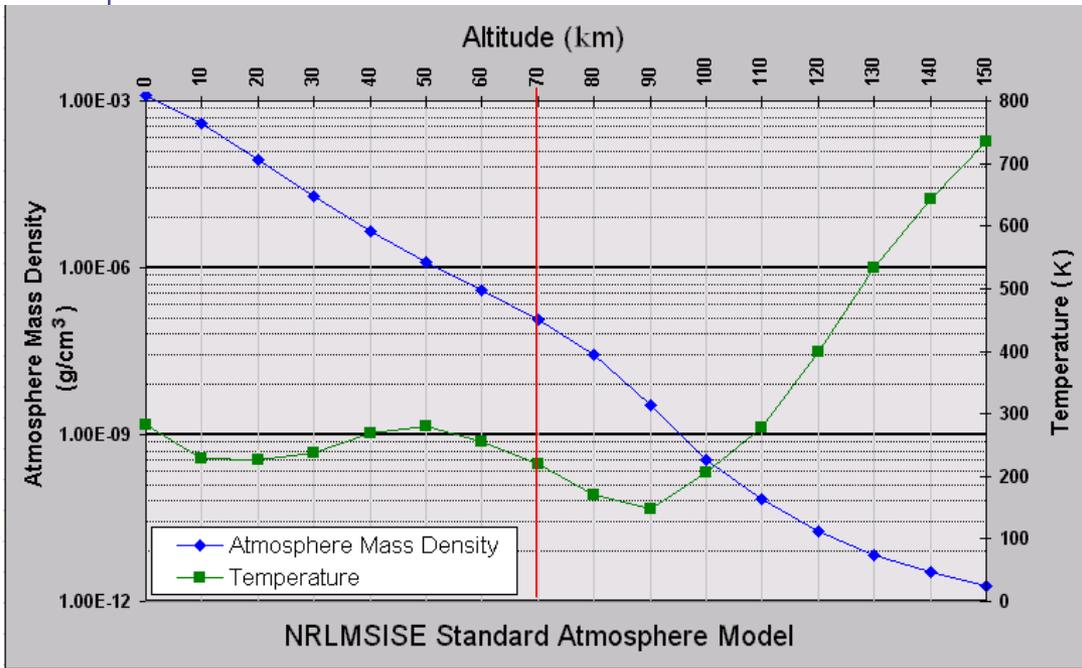
- a main series of layers made from the part of the atmospheric shells **between the two cones** (this is the part where the scoring takes place)
- two series of side layers made from the part of the atmospheric shells **between one of the two cones and the third one**

*These additional layers are needed to take into account the primary and secondary particles which don't come from the vertical direction but can anyway reach the region of interest.*

A file **atmlocmat.cards** will contain additional material assignments to be included in the input.



# Atmospheric model: density (1)



The atmosphere can be roughly characterized as the region from sea level to about 1000 km altitude around the globe, where neutral gases can be detected.

**Below 50 km the atmosphere can be assumed to be homogeneously mixed and can be treated as a perfect gas. Above 80 km the hydrostatic equilibrium gradually breaks down as diffusion and vertical transport become important. The FLUKA atmospheric model, below 70 km, is in the perfect gas region.**

# Atmospheric model: density (2)

This Table shows the U.S. Standard Atmosphere depth vs altitude and vs FLUKA atmospheric layer.

**100 layers** from 0 to 70 km above sea level.

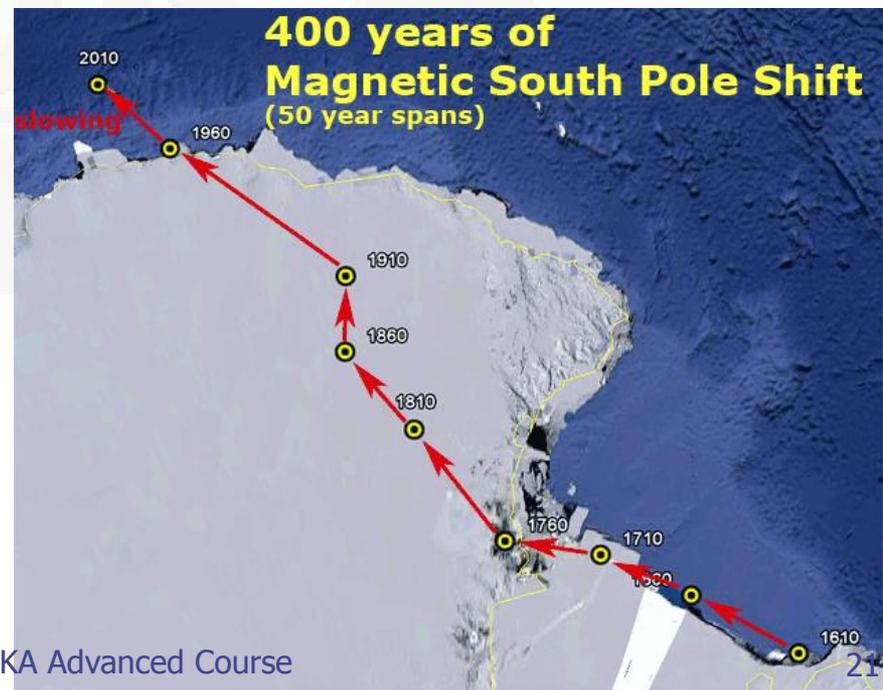
They are described as **100** different FLUKA **regions**, corresponding to 100 different materials (air, with different densities)

FLUKA region	km from s.l.	US St. Atm. Depth (g/cm <sup>2</sup> )	FLUKA region	km from s.l.	US St. Atm. Depth (g/cm <sup>2</sup> )	FLUKA region	km from s.l.	US St. Atm. Depth (g/cm <sup>2</sup> )
1.0	70.0	0.092	35.0	31.6	9.367	69.0	10.7	242.777
2.0	68.5	0.108	36.0	30.8	10.540	70.0	10.2	260.107
3.0	67.1	0.126	37.0	30.0	11.849	71.0	9.8	278.093
4.0	65.6	0.146	38.0	29.2	13.309	72.0	9.4	296.729
5.0	64.2	0.170	39.0	28.4	14.937	73.0	8.9	316.007
6.0	62.8	0.198	40.0	27.7	16.748	74.0	8.5	335.921
7.0	61.5	0.230	41.0	26.9	18.763	75.0	8.1	356.460
8.0	60.1	0.266	42.0	26.2	21.004	76.0	7.7	377.615
9.0	58.8	0.308	43.0	25.5	23.492	77.0	7.3	399.374
10.0	57.5	0.356	44.0	24.8	26.255	78.0	6.9	421.727
11.0	56.2	0.411	45.0	24.1	29.290	79.0	6.6	444.661
12.0	55.0	0.474	46.0	23.4	32.613	80.0	6.2	468.163
13.0	53.8	0.546	47.0	22.7	36.244	81.0	5.8	492.219
14.0	52.5	0.628	48.0	22.1	40.205	82.0	5.5	516.815
15.0	51.4	0.722	49.0	21.4	44.516	83.0	5.1	541.936
16.0	50.2	0.828	50.0	20.8	49.201	84.0	4.8	567.566
17.0	49.1	0.950	51.0	20.2	54.283	85.0	4.4	593.691
18.0	47.9	1.088	52.0	19.6	59.785	86.0	4.1	620.295
19.0	46.8	1.245	53.0	19.0	65.733	87.0	3.8	647.359
20.0	45.7	1.423	54.0	18.4	72.152	88.0	3.4	674.869
21.0	44.7	1.625	55.0	17.8	79.068	89.0	3.1	702.807
22.0	43.6	1.854	56.0	17.2	86.506	90.0	2.8	731.155
23.0	42.6	2.112	57.0	16.7	94.493	91.0	2.5	759.898
24.0	41.6	2.404	58.0	16.1	103.057	92.0	2.2	789.016
25.0	40.6	2.734	59.0	15.6	112.224	93.0	1.9	818.493
26.0	39.6	3.106	60.0	15.0	122.023	94.0	1.6	848.311
27.0	38.7	3.525	61.0	14.5	132.482	95.0	1.3	878.453
28.0	37.7	3.996	62.0	14.0	143.628	96.0	1.1	908.900
29.0	36.8	4.526	63.0	13.5	155.489	97.0	0.8	939.636
30.0	35.9	5.121	64.0	13.0	168.094	98.0	0.5	970.643
31.0	35.0	5.789	65.0	12.5	181.471	99.0	0.3	1001.903
32.0	34.1	6.538	66.0	12.0	195.646	100.0	0.0	1033.400
33.0	33.3	7.378	67.0	11.6	210.649			
34.0	32.4	8.317	68.0	11.1	226.507			

# Geomagnetic field (1)

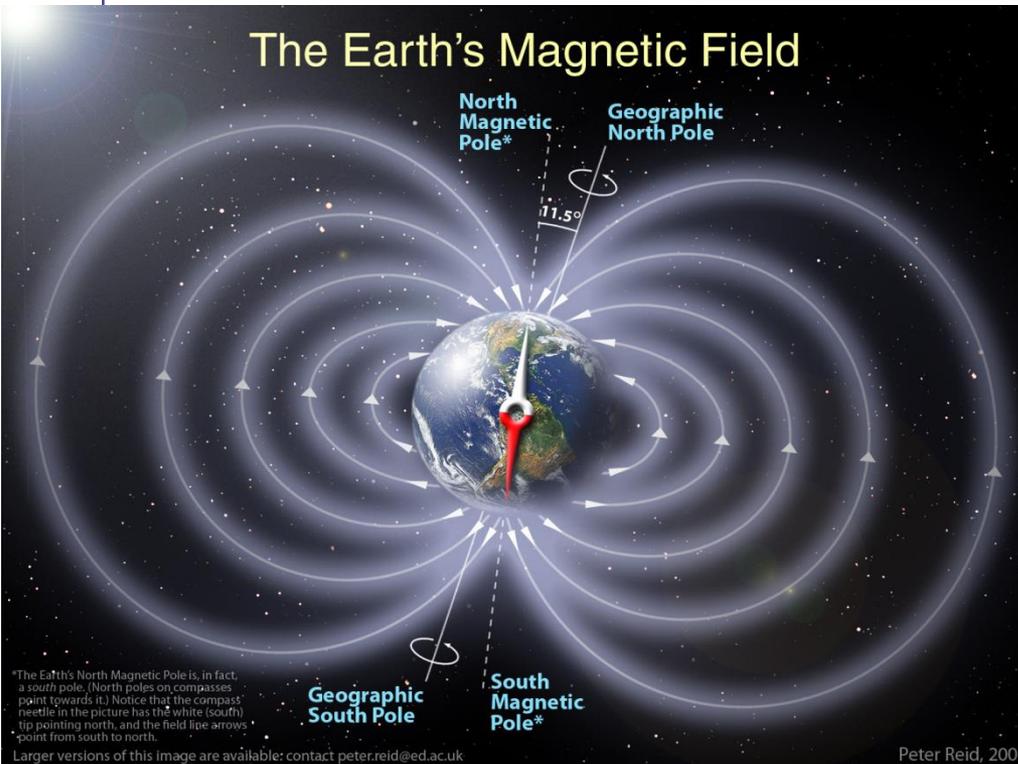


In the last 50 years measurements of the geomagnetic field configuration have been performed regularly with increasing precision, revealing a yearly weakening of the field intensity of 0.07% and a westward drift of  $\sim 0.2$  degrees per year over the Earth's surface.



# Geomagnetic field (2)

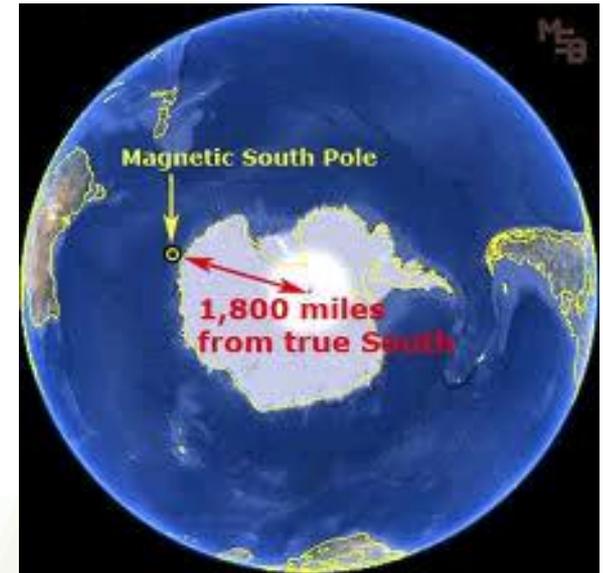
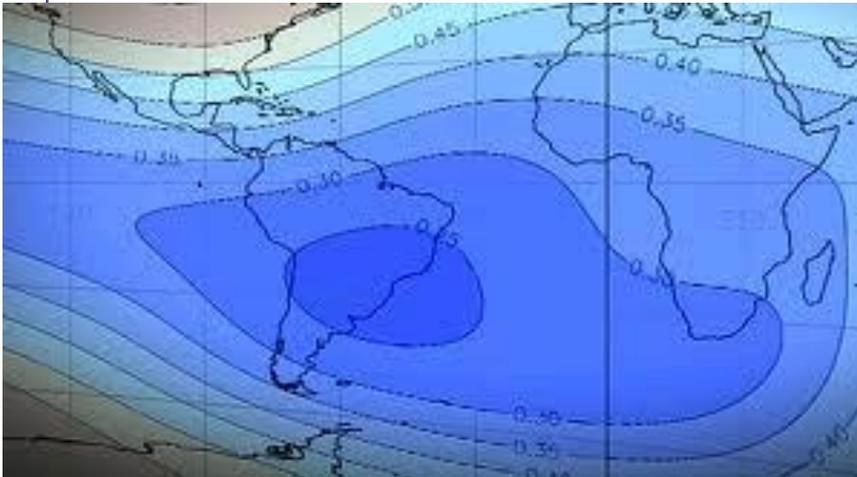
## The Earth's Magnetic Field



This field can be described, to first order, as a **magnetic dipole tilted** with respect to the rotation axis by  $\sim 11.5$  degrees, and **displaced** by  $\sim 400$  km with respect to the Earth's center and with a magnetic moment  $M = 8.1 \times 10^{25} \text{ G cm}^3$  ( $8.1 \times 10^{22} \text{ A m}^2$ ). The dipole orientation is such that the magnetic South pole is located near the geographic North pole, in Greenland, at a latitude of  $75^\circ \text{ N}$  and a longitude of  $291^\circ$ .

# Geomagnetic field (3)

The magnetic North pole is instead near the geographic South pole, on the border of the Antarctica →



The intensity at the Earth's surface (above) varies from a maximum of  $\sim 6 \times 10^{-5}$  T near the magnetic poles to a minimum of  $\sim 2 \times 10^{-5}$  T in the region of the South Atlantic Anomaly (SAA), between Brazil and South Africa. The complex behavior of the equipotential field lines is mainly a consequence of the offset and tilt.

***Please note that the overall intensity is not the most important parameter, rather the orientation is what matters most. Indeed the (magnetic) polar regions despite having the largest intensity are the least effective in shielding from the GCR's***

# Geomagnetic field (4)

In FLUKA the geomagnetic field is **taken into account in two different stages:**

- 1) **Effect of geomagnetic cutoff which modulates the primary spectrum:** at a given location (point of first interaction of primary particles) and for a given direction, a **threshold in magnetic rigidity** exists. The closer the injection point is to the geomagnetic equator, the higher will be the vertical rigidity threshold. The standard possibilities offered to the user are:
  - a) For **"local"** geometries: evaluate the geomagnetic cutoff making use of a dipolar field centered with respect to the centre of the Earth, adapted to give the "correct" vertical rigidity cutoff for the geographic location under examination
  - b) For **"global"** geometries:
    - i. evaluate the geomagnetic cutoff making use of a dipolar field centered with respect to the centre of the Earth
    - ii. evaluate the geomagnetic cutoff making use of a properly offset dipolar field
    - iii. as ii. down to some altitude and then using IGRF11 for the "exact" magnetic field
- 2) **The local geomagnetic field can be taken into account during shower development in the atmosphere.** The field is automatically provided by the default **MAGFLD** FLUKA user routine, in accordance to the option selected in the **GCR-SPE** card. *For all problems, provided the coordinate system is consistently used (that is, geomagnetic coordinates for the centered dipolar field, geographic ones for the offset dipole or multipolar field) there is no need to provide any orientation or intensity information about the field.*

# Two commands for Cosmic Rays

## GCR-SPE

Initializes Galactic Cosmic Ray or Solar Particle Event calculations

## SPECSOUR

defines one of the following special sources:

- ❑ Galactic Cosmic Rays  
(SDUM = GCR-IONF, GCR-SPEC, GCR-ALLF)
- ❑ Solar Particle Event  
(SDUM = SPE-SPEC, SPE-2003, SPE-2005)

The usual scoring options (**USRBDX**, **USRTRACK...**) can be used to define detectors to calculate the fluence of different radiation fields. *For estimators in the "standard" atmospheric layers an automatic normalization factor is applied → results are directly in p/cm<sup>2</sup>/s, check the output to know whether this is the case!!*

# Cosmic ray tools (1)

- ❑ A number of tools and packages have been developed for the FLUKA environment to simulate the production of secondary particles by primary cosmic rays interacting with the Earth's atmosphere. These tools, in different stand-alone versions, have already been successfully used for fundamental physics research
- ❑ The set of FLUKA tools for cosmic ray simulation includes a set of core routines to manage **event generation**, **geomagnetic effects** and **particle scoring**, a **standalone program**, and a number of **stand-alone data files**

## The standalone program

- ❑ **atmloc\_2011.f**: prepares the description of the local atmosphere geometry with the atmospheric shells initialized by option **GCR-SPE**. This geometry includes only a slice of the Earth geometry, centered around the geomagnetic latitude input by the user

# Cosmic ray tools (2)

## The files

- **atmomat.cards**: contains the material definitions for the density profile of the US Standard Atmosphere. These cards must be inserted (or the file included with the **#include** directive) into the FLUKA input file.
- **atmogeocards**: contains an example of a 3D geometrical description of the Earth atmosphere, generated in accordance with the previous data cards (and corresponding density profile). This geometry includes the **whole Earth**
- **<iz>phi<MV>.spc**: GCR All-Particle-Spectra for the **iz<sup>th</sup>** ion species (**iz** = 1,...,28), modulated for the solar activity corresponding to a **Phi** parameter **<MV>** MegaVolt. **Phi=500 MV** roughly corresponds to solar minimum, while **Phi=1400 MV** roughly corresponds to solar maximum
- **allnucok.dat**: GCR All-Nucleon Spectra
- **sep20jan2005.spc**: spectra for the Solar Particle Event of Jan 20th, 2005
- **sep28oct2003.spc**: spectra for the Solar Particle Event of Oct 28th, 2008

# From the manual: SPECSOUR/GCR-IONF (1)

- ❑ SDUM = GCR-IONF: All-particle flux  
The particle composition of the flux can be modified by choosing the minimum and maximum atomic number ( $1 \leq Z \leq 28$ ).
- ❑ The spectrum components have been produced by a code for various modulation parameters and written on '.spc' files ( $Z + \langle \text{PhiMV} \rangle + .\text{spc}$ ).
- ❑ It is possible to give an energy interval and to choose a starting radius\* (radius of the emission sphere in case of spherical geometry) or starting height (the emission height in case of flat geometry).
- ❑ It is possible to activate the geomagnetic cutoff (WHAT(7) in SPECSOUR) and to input optionally the vertical cutoff value at the central latitude.
- ❑ Ions are treated like real ions or can be split. The optimized value for spectral index for sampling (below transition energy) is  $\gamma = 1.75$  (WHAT(5)).
- ❑ Above transition energy, the spectrum will be assumed for sampling purposes to have a  $1/E$  shape. Obviously weights are properly adjusted

*\*When using the offset dipole, or the offset dipole+IGRF11, the emission radius is centered on the dipole center*

# Summary from the manual: SPECSOUR/GCR-IONF (2)

## SDUM = GCR-IONF: All-particle flux

- ❑ **WHAT (1)** =  $Z_{\max} + 100 * Z_{\min}$  ( $Z_{\min} = 1$  if none is defined)
- ❑ **WHAT (2)** = Starting radius\* or starting height (cm)
- ❑ **WHAT (3)** = Minimum energy
- ❑ **WHAT (4)** = Maximum energy  
If max. and minimum energy differ by  $< 5\%$  then a fixed energy (= Max) is sampled
- ❑ **WHAT (5)** = Spectral index for sampling (below transition energy)
- ❑ **WHAT (6)** = Transition energy for sampling (above it, sample from  $1/E$ )

## Continuation card (**SDUM = "&"**)

- ❑ **WHAT (7)** = 0: no geomagnetic cutoff;  
= 1: geomagnetic cutoff is requested (centered dipole)  
= 2: the vertical geomagnetic cutoff is read from WHAT(2)  
= 3: geomagnetic cutoff is requested (offset dipole, maybe + IGRF11)
- ❑ **WHAT (8)** = vertical geomagnetic cutoff at central latitude for WHAT(1) = 2
- ❑ **WHAT (9)** = number of energy point in the spectra. Default: 50
- ❑ **WHAT(10)** = if  $> 0$  vertical run (*for testing purposes only!*)
- ❑ **WHAT(11)** = if  $> 0$  probabilities  $1/(2xZ)$  are used for the various ions (1 for  $Z = 1$ ) *deprecated!*
- ❑ **WHAT(12)** =  $< 0$ : ions are split  
=  $> 0$ : ions are treated like real ions

*\*When using the offset dipole, or the offset dipole+IGRF11, the emission radius is centered on the dipole center*

# Summary from the manual: SPECSOUR/SPE-xxxx

**SDUM = SPE-SPEC, SPE-2003 or SPE-2005:** Solar Particle Event.

**SPE-SPEC, SPE-2005** → spectrum is read from file `sep20jan2005.spc`

**SPE-2003** → spectrum is read from file `sep28oct2003.spc`

The WHATs are the same as for **SDUM=GCR-IONF**

A few caveats:

- Contrary to GCR spectra, SPE spectra are often not well known particularly at the highest energies (most data come from satellites which have inherent limitations)
- Ions ( $Z > 1$ ) spectra in the provided files are educated guesses
- Every SPE is different in spectrum/intensity/duration
- Real SPE's are not isotropic, but rather reflect the geometry of the solar wind impact on earth
- Spectra can change during a SPE event, those provide are integrated (averaged) ones for those two events

# From the manual: SPECSOUR/GCR-ALLF (1)

- ❑ **SDUM = GCR-ALLF: All-nucleon flux**  
Three different options (average, maximum and minimum flux) are available.
- ❑ The program reads fluxes from a file named "**allnucok.dat**" in which are given the **total energy** (GeV), the **fluxes** (E.dN/dE) and the **neutron/proton ratios**.
- ❑ It is possible to give an energy interval and to choose:
  - ❑ a starting radius (radius of the emission sphere in case of spherical geometry)\*  
or
  - ❑ a starting height (the emission height in case of flat geometry).
- ❑ It is possible to activate the vertical geomagnetic cutoff and to give the cutoff value at the central latitude, otherwise the geomagnetic cutoff will be not taken into account.
- ❑ Ions are treated as separate nucleons, or as alphas and protons.

*\*When using the offset dipole or the offset dipole+IGRF11 the emission radius is centered on the dipole center*

# From the manual: SPECSOUR/GCR-ALLF (2)

**SDUM = GCR-ALLF:** All-nucleon flux

Most **WHAT's** are common to **GCR-IONF**, only those which are different are listed below:

- ❑ **WHAT (1) = 1:** central value for the all-nucleon fit;  
= 2: lower limit;  
= 3: upper limit;

Continuation card (SDUM = "&"):

- ❑ **WHAT(12) =< 0:** nucleons are transported separately  
> 0: transport as many alphas as can be built by neutrons, and the remaining protons

# Summary from the manual: GCR-SPE, Notes

Notes to GCR-SPE :

- 1) Cosmic ray calculations, initialized by GCR-SPE are defined by means of command SPECSOUR and a number of auxiliary programs.
- 2) The cards for the geometry description of the atmospheric shells must be prepared using the auxiliary programs and data cards in the directory \$FLUPRO/gcrttools. Program atmloc\_2011.f writes a file atmloc.geo, containing the geometry input to be inserted into the FLUKA input file (or to be read by setting WHAT(3) in the GEOBEGIN card), a file atmlocmat.cards containing the extra material assignments, and a file atmloc.sur containing auxiliary data and the scoring areas.

The user shall rename the file atmloc.sur to <xxxxxxx>.sur, where <xxxxxxx> is an identifier of exactly 7 characters which must appear also in the input spectra file names: the spectra must have the names <zz><xxxxxxx>.spc, where <zz> is the atomic number of the ion.

The example spectra distributed with FLUKA come with two identifiers: <zz>phi0465 for solar minimum and <zz>phi1440 for solar maximum, and with zz=01-28.

# Summary from the manual: GCR-SPE (2)

## GCR-SPE: Input parameters initialization

### SDUM=" "

- **WHAT (1)** =  $i0 + 1000 * i1$ 
  - $i0 = 0$ : naïve centered dipole field, geomagnetic coordinates
  - $= 1$ : exact multipolar expansion field (now based on IGRF11), geographic coordinates
  - $= 4$ : offset dipole field, geographic coordinates
- $i1$  = number of atmospheric shells (51,101,201)
- **WHAT (2,3,4,6)** = not used
- **WHAT (5)** >0: Planet equatorial magnetic field (T)

### SDUM = "DIPCOORD" (for offset dipole/IGRF11 only)

- **WHAT (1)** = dipole center position radius (km)
- **WHAT (2)** = dipole center position latitude (deg)
- **WHAT (3)** = dipole center position longitude (deg)
- **WHAT (4)** = geomagnetic North Pole latitude (deg)
- **WHAT (5)** = geomagnetic North Pole longitude (deg)
- **WHAT (6)** = year (for IGRF11 only), it should be in the 1900-2015 range, def. 2010.0

# Example of input data cards

Examples of user data cards to run a FLUKA cosmic ray problem are provided with the code. Eg, an example refers to the simulation at geographical coordinates of 36.0 degrees North Latitude and 140.0 degrees East Longitude (Tsukuba, Japan), using the solar modulation of Dec. 23rd 1995:  $\mu^+$  and  $\mu^-$  fluxes at different heights in the atmosphere are then scored.

**Example available in:**

**\$FLUPRO/gcrtools/gcrexamples/AllParticleExample**

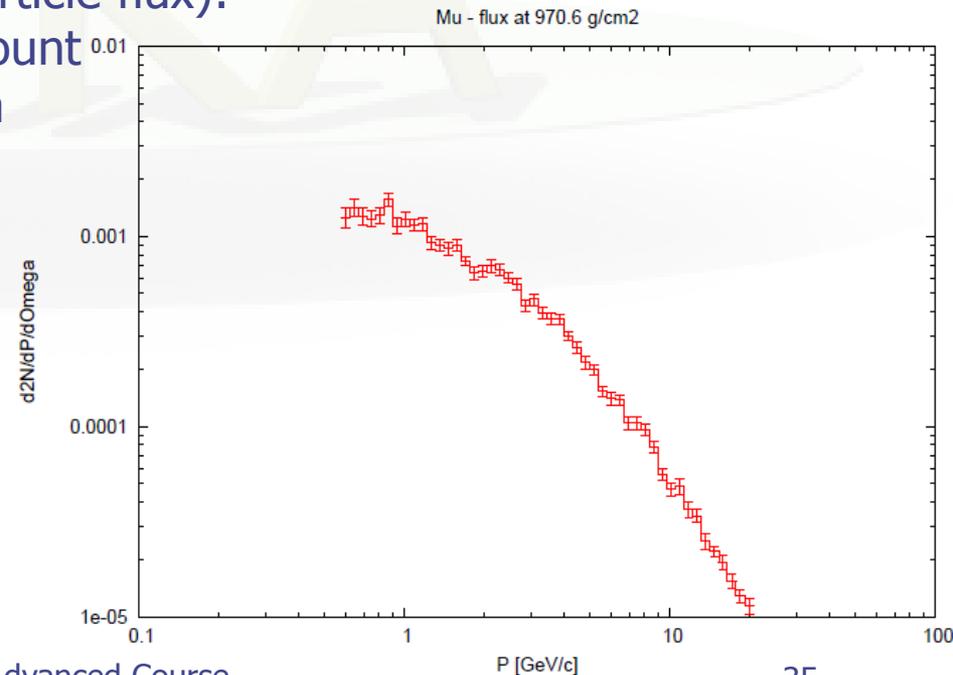
**SPECSOUR** and **SDUM = GCR-IONF** (all particle flux):

Z of ions of spectra to be taken into account

- information for sampling energy spectra
- geomagnetic cutoff
- starting radius

**GCR-SPE:** Initializes Galactic Cosmic Ray or Solar Particle Event calculations :

- geomagnetic field
- spectra files name



# Example:

## **GCR-SPE** initializes Galactic Cosmic Ray or Solar Particle Event calculations

```
GCR-SPE 10100.
```

```
phi0465
```

What(1): various options for the magnetic field (100 atm shells, naïve dipolar field)

SDUM: name of spectra files (read spectra from zzphi0465.spc)

## **SPECSOUR with SDUM GCR-SPE** (SPECSOUR doesn't only define GCR spectra – see before) **calls special GCR source**

```
SPECSOUR 28.0 6.449D+08 0.3 30000. 1.75 500. GCR-IONF  
SPECSOUR 2.0 11.4 &
```

### First card:

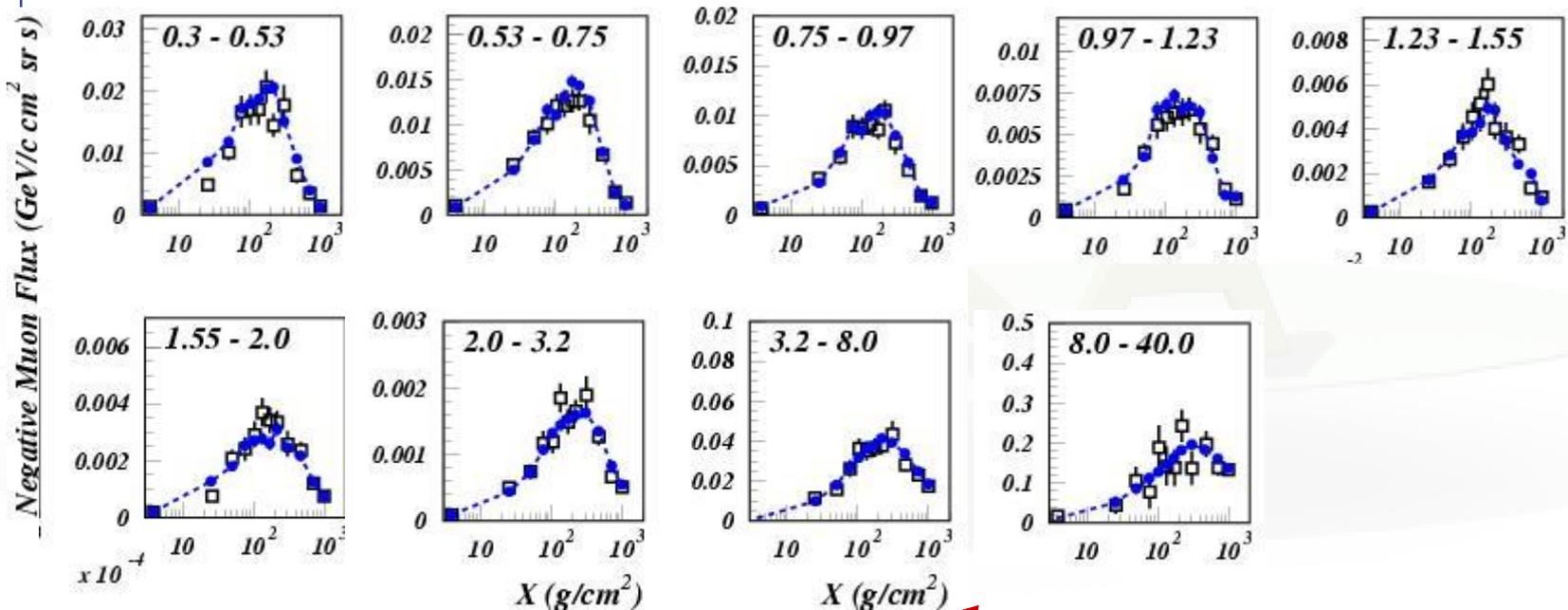
(1) Z range, (2) Inj.Radius, (3) Emin, (4) Emax, (5) Sampling index, (6) Transition energy

### Second card:

(7) cutoff?, (8) cutoff, (9)(# energy points), (10)(vertical run), (11)(ion probabilities), (12) split

# Muon benchmark: CAPRICE

**Momentum bin (GeV/c)** *Open symbols: CAPRICE data*  
*Full symbols: FLUKA*



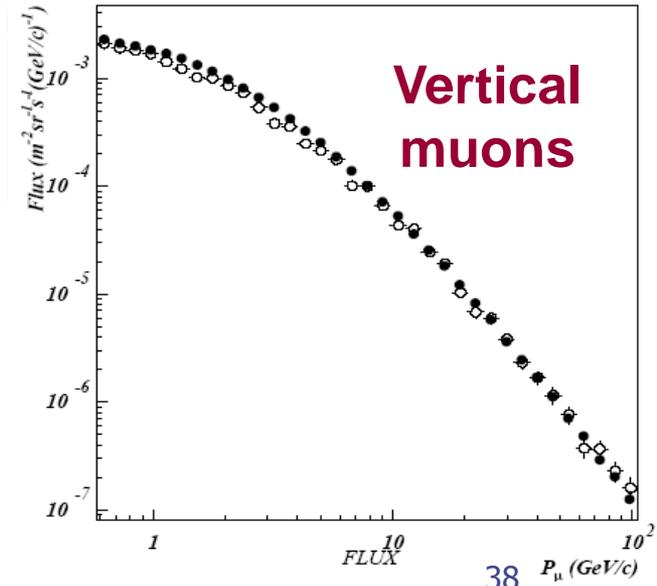
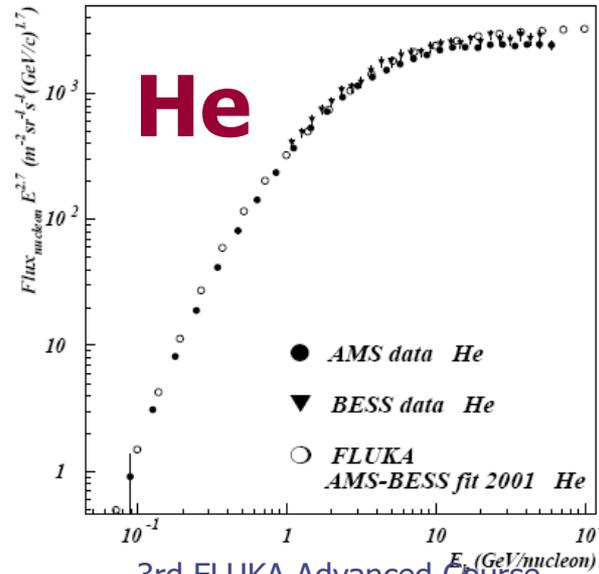
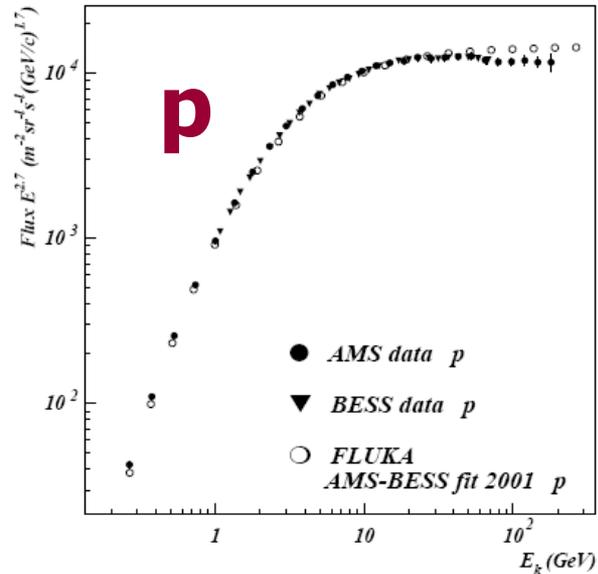
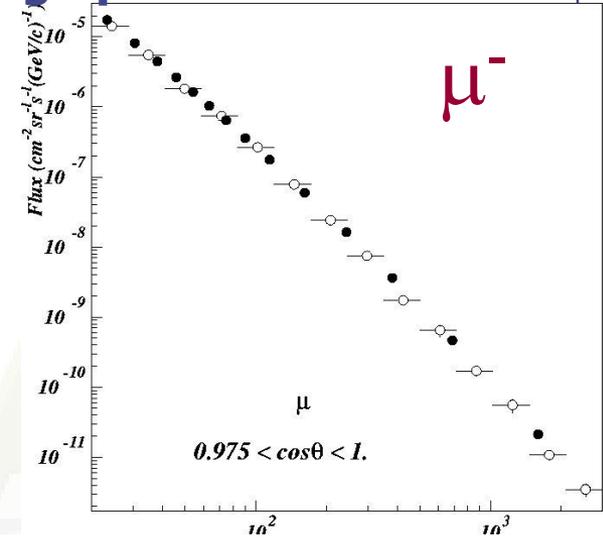
**Primary spectrum normalization  $\sim$ AMS-BESS**  
***Astrop. Phys., Vol. 17, No. 4 (2002) p. 477***

***Atmospheric  
thickness  
( $\text{g/cm}^2$ )***

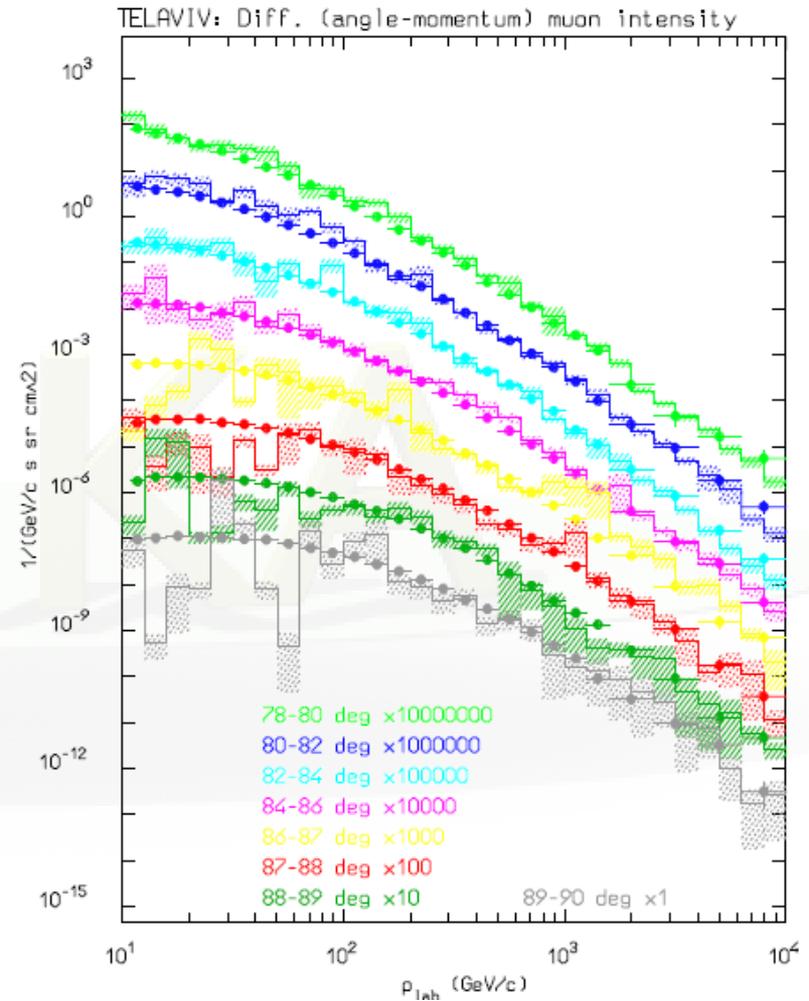
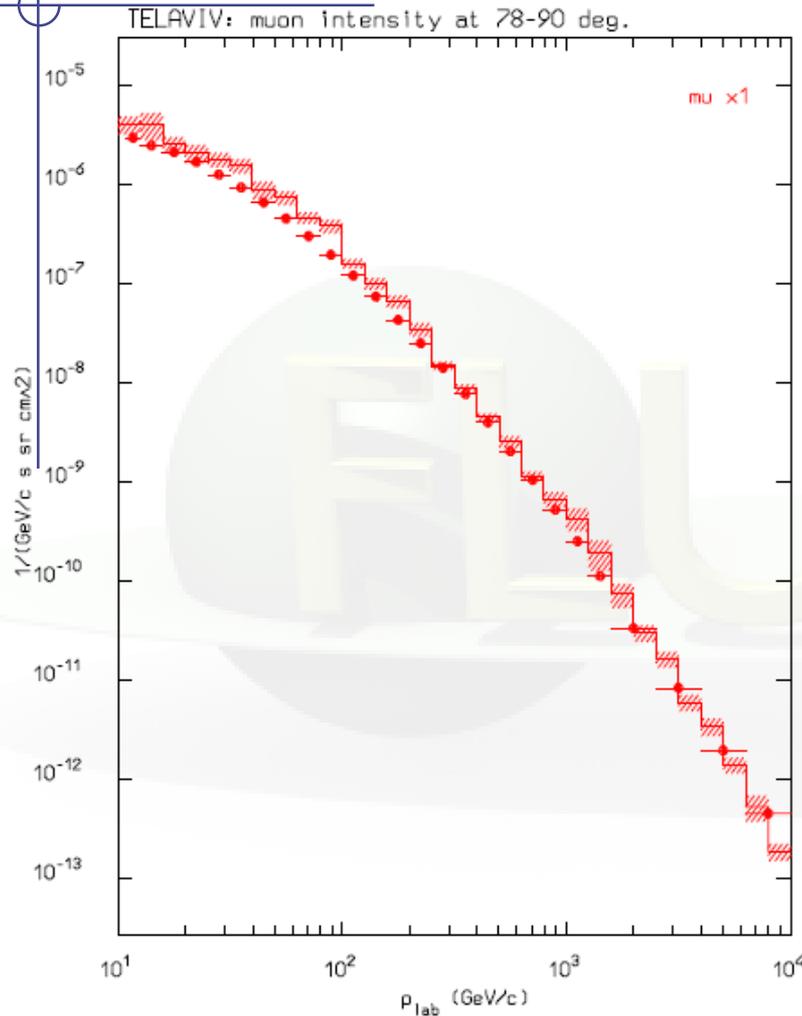
# BESS spectrometer

Balloon-borne Experiment with Superconducting Spectrometer

The BESS spectrometer has collected data at different cut-offs, altitudes, solar modulation

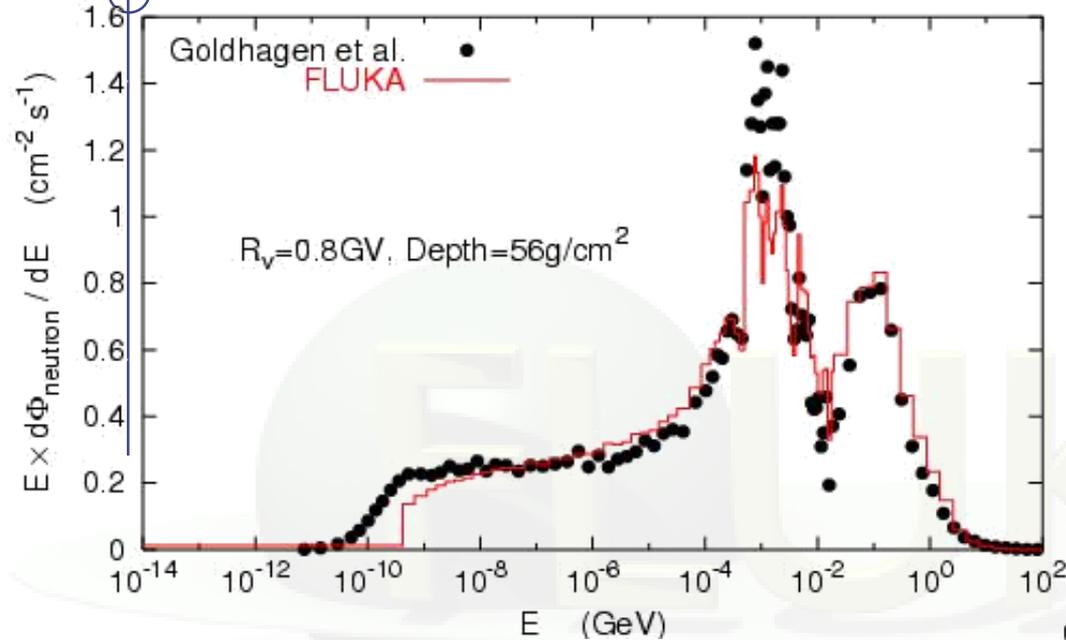


# Benchmark: Earth Surface



Angle integrated (78-90deg, left) and double differential muon fluxes, measured close to the horizontal line TelAviv. Data from O.C. Allkofer et al. NPB 259,1, (1985).

# Neutrons on the ER-2 plane at 21 km



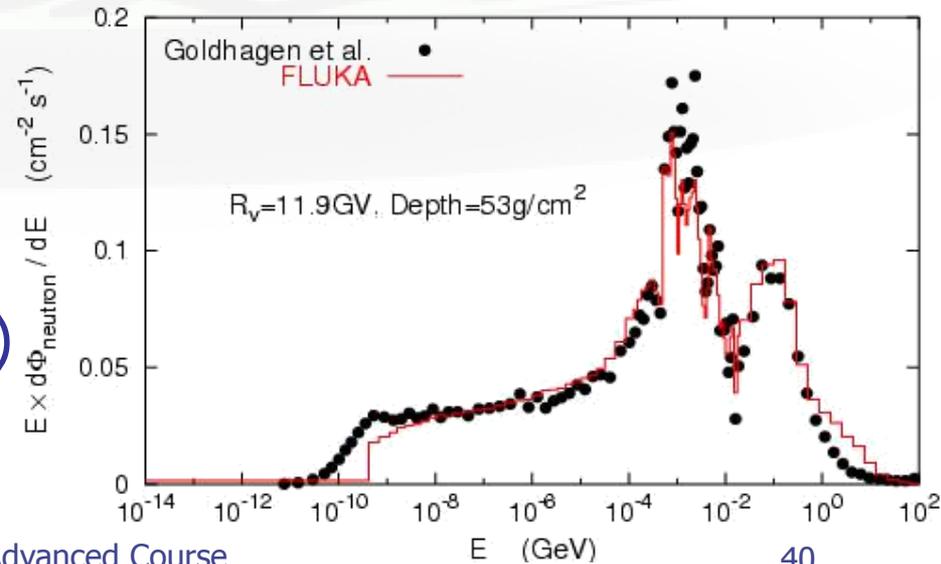
## Measurements:

Goldhagen et al.,  
NIM A476, 42 (2002)

Note one order of magnitude  
difference depending on  
latitude

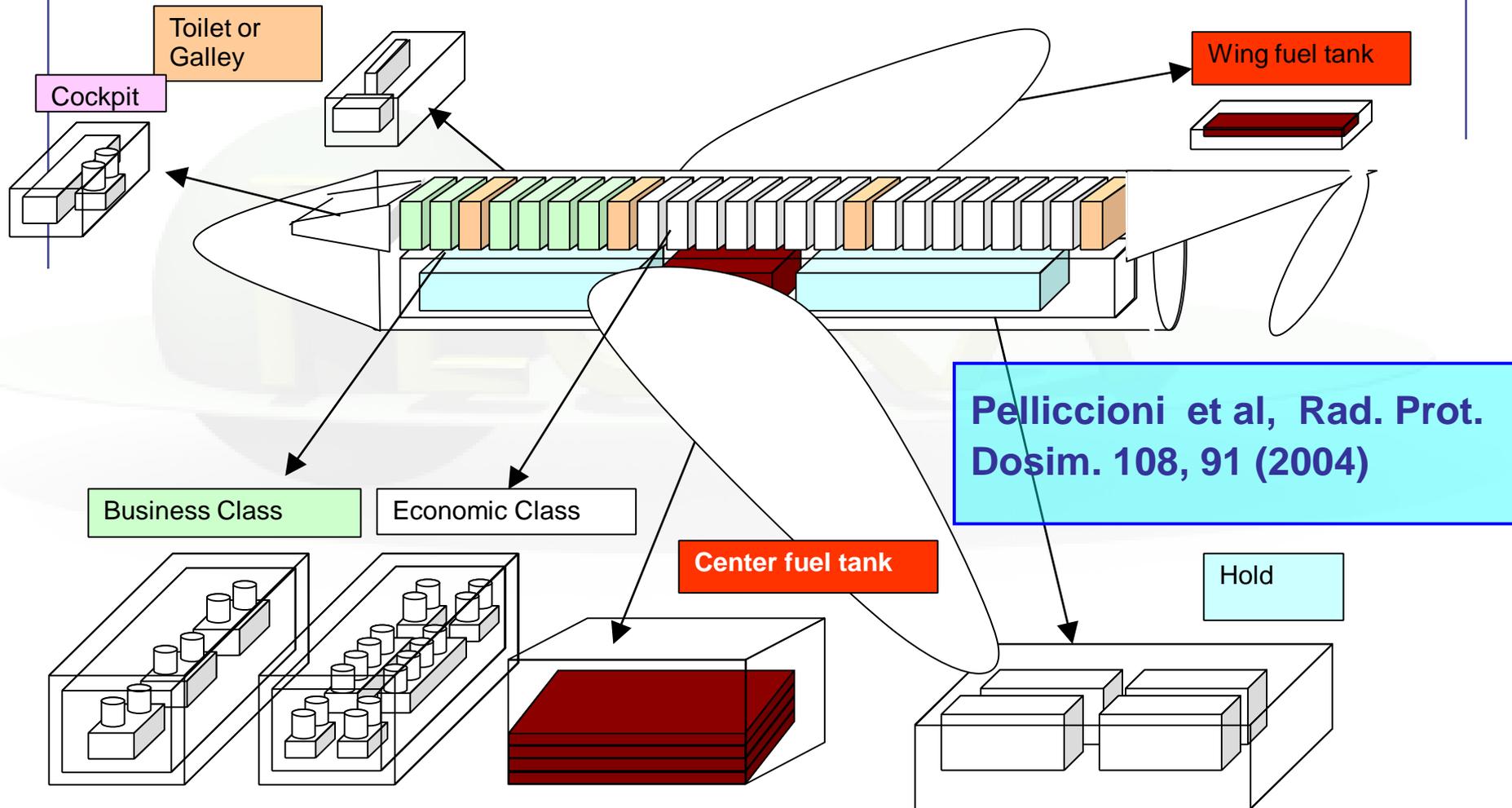
## FLUKA calculations:

Roesler et al.,  
Rad. Prot. Dosim. 98, 367 (2002)



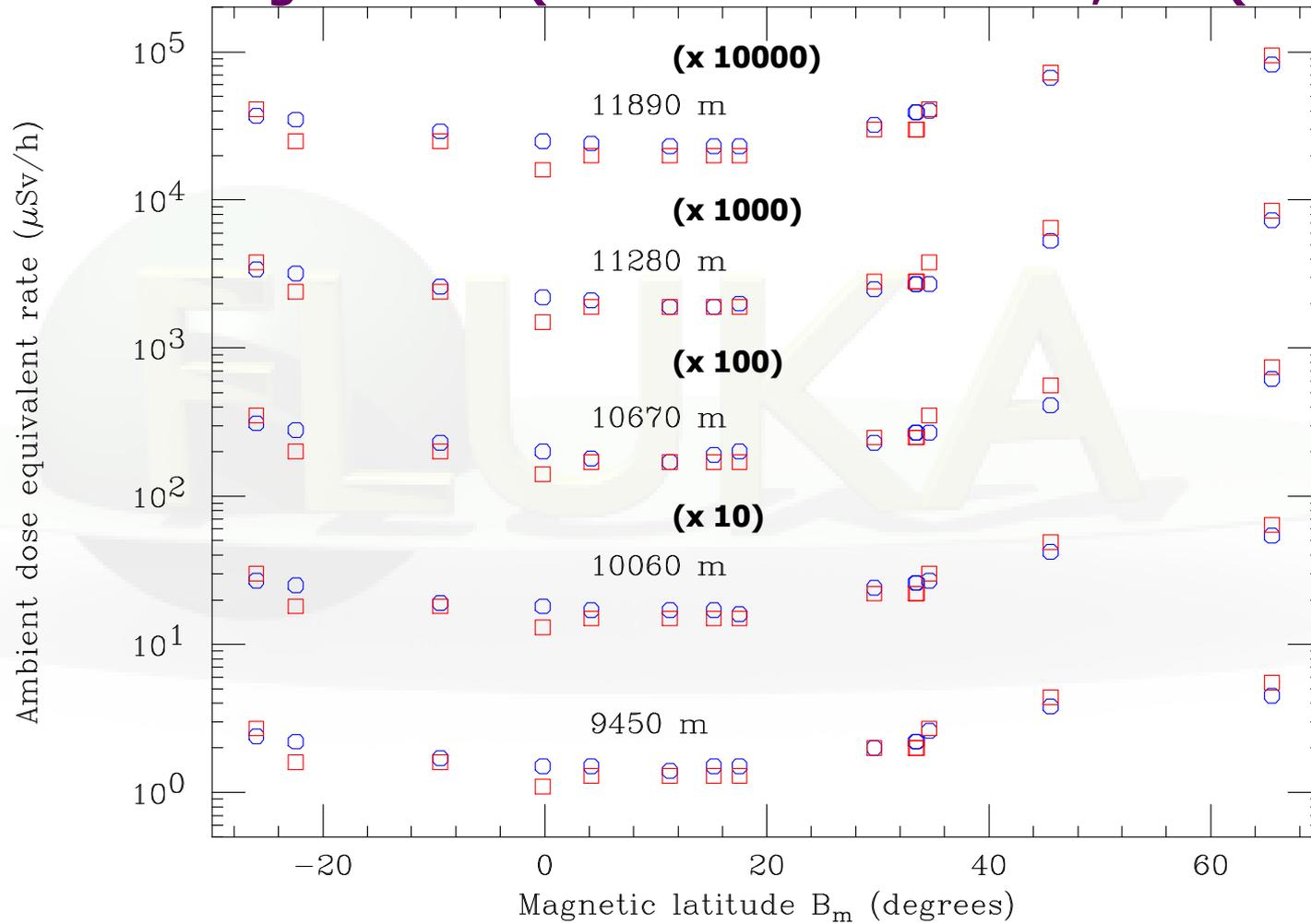
# Doses to aircrew and passengers

## AIRBUS 340



# Doses to aircrew and passengers

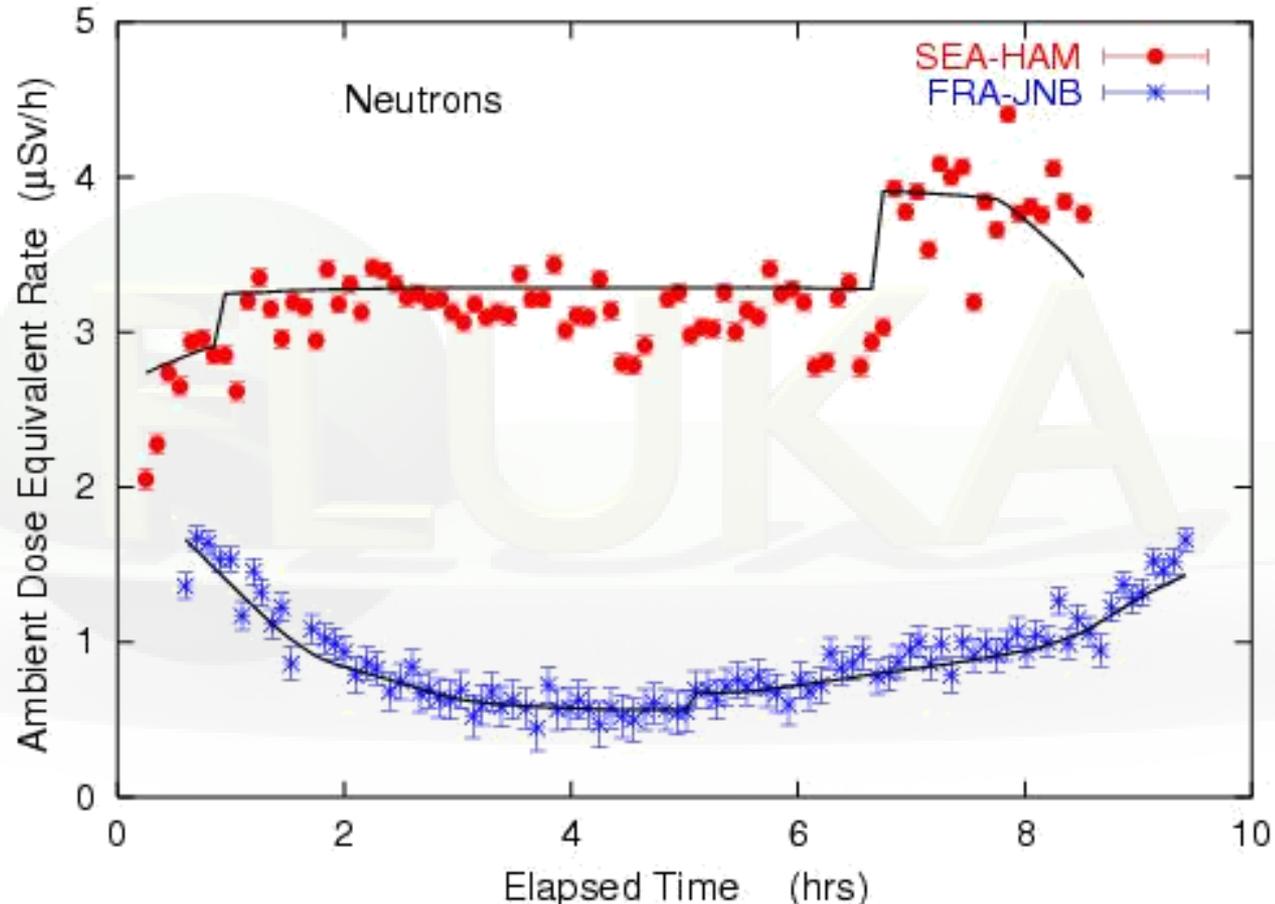
Commercial flight doses: (Pelliccioni et al. RPD93, 101 (2001))



Simulated (FLUKA, red) and measured (blue, NIMA422, 621, 1999) ambient dose equivalent for various altitudes (scaled by one decade) and geomagnetic cut-off's

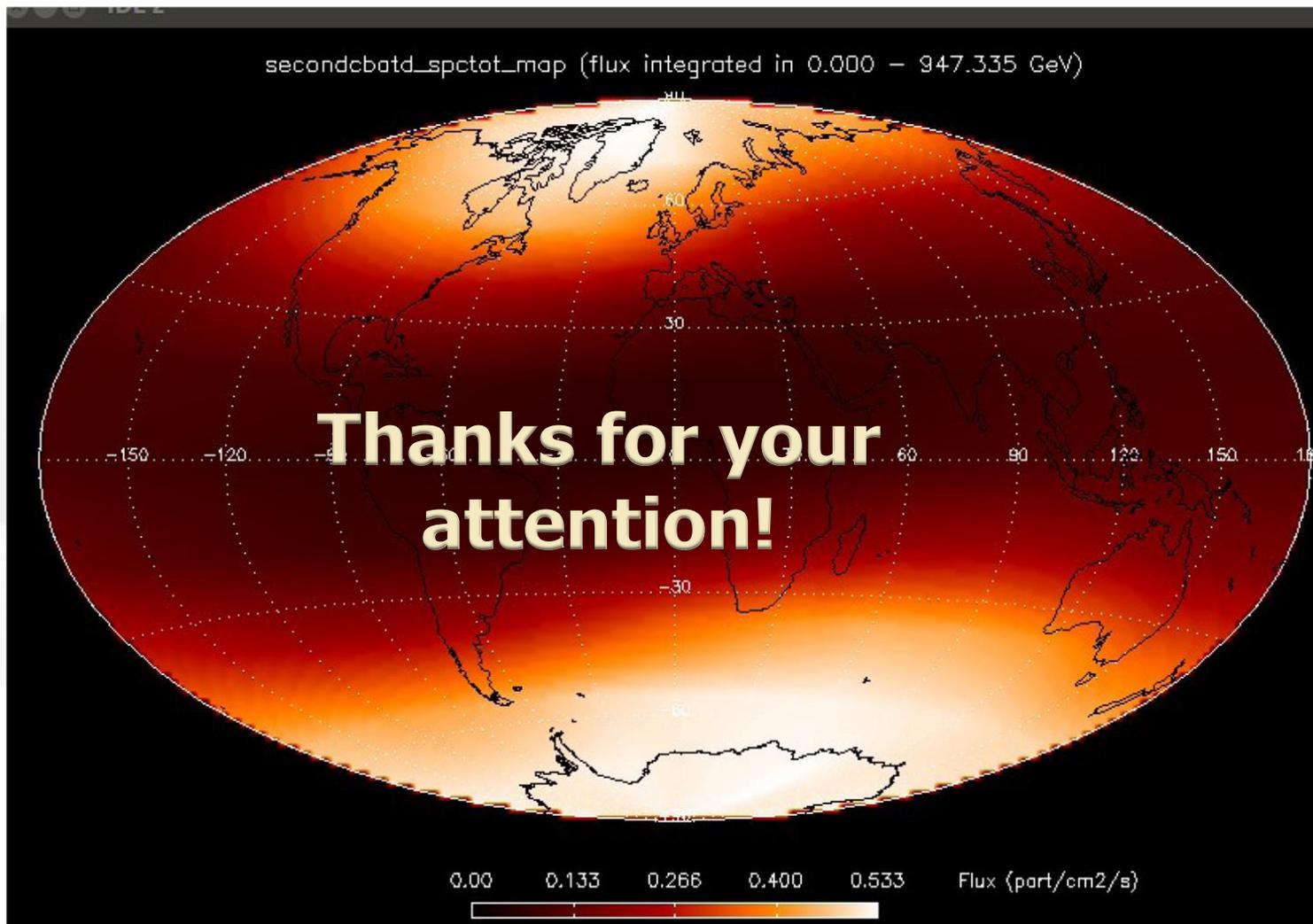
# Doses to aircrew and passengers

Roesler et al., Rad. Prot. Dosim. 98, 367 (2002)



Ambient dose equivalent from neutrons at solar maximum on commercial flights from Seattle to Hamburg and from Frankfurt to Johannesburg. **Solid lines: FLUKA simulation**

# The neutron albedo at 400 km altitude



**Offset dipole + IGRF11 with parameter adjusted at year 2010**