

#### Introduction to Particle Physics

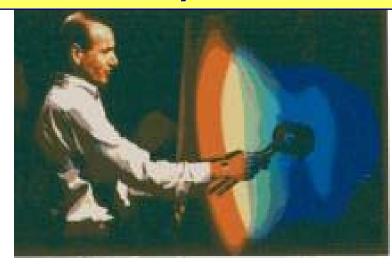
(for non physics students)

#### 4. UNIFIED UNIVERSE

(no strings attached)



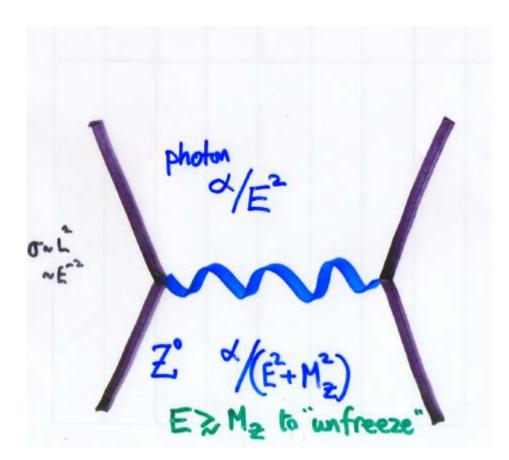
PROFESSOR FRANK CLOSE
EXETER COLLEGE
UNIVERSITY OF OXFORD



FORCES Summary

(remember that waves -> particles)

NAME	action	CARRIER
Gravity	ground	graviton?
Electromo	solids stops us falling to centre of Earth	byoton (8)
Weak	β-radioactivity P→He in Sun	W+ W- Z0
Strong	quatks glued inside p,n  P,n in nuclei	gluons (g) 8 different
Only the	weak force carriers	have MASSes
My ~	80 GeV/c2	
M.~	91 GeV/c2	



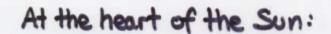
#### Feynman rules:

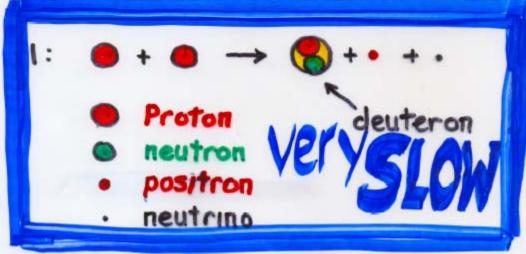
If energy E flows through the transmitted "virtual" particle (photon; Z) it costs 1/(E^2+M^2)



If E >> M the cost is 1/E^2....like the case of the photon

Only appears weak at low energy. Unified at high energy







(2:  $\bullet$  +  $\bullet$   $\rightarrow$   $\bullet$  (3He)?

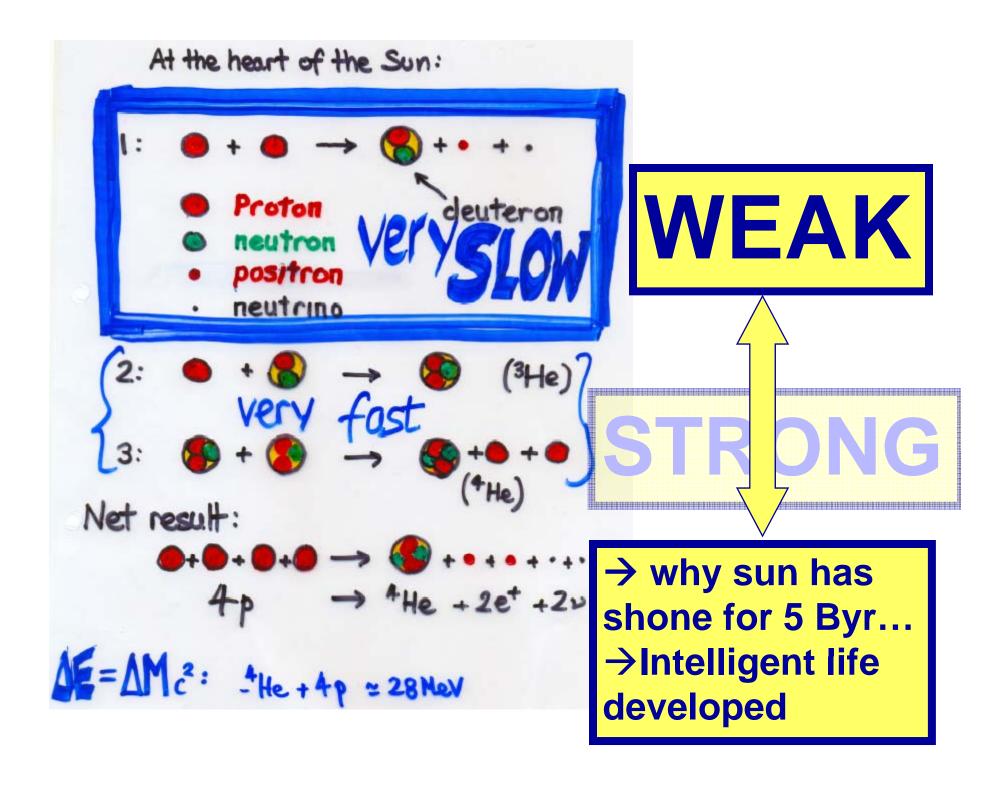
Very fast

3:  $\bullet$  +  $\bullet$   $\rightarrow$   $\bullet$  +  $\bullet$  +  $\bullet$ 

## STRONG

Net result:

M= ΔMc2: - He + 4p = 28 Nov

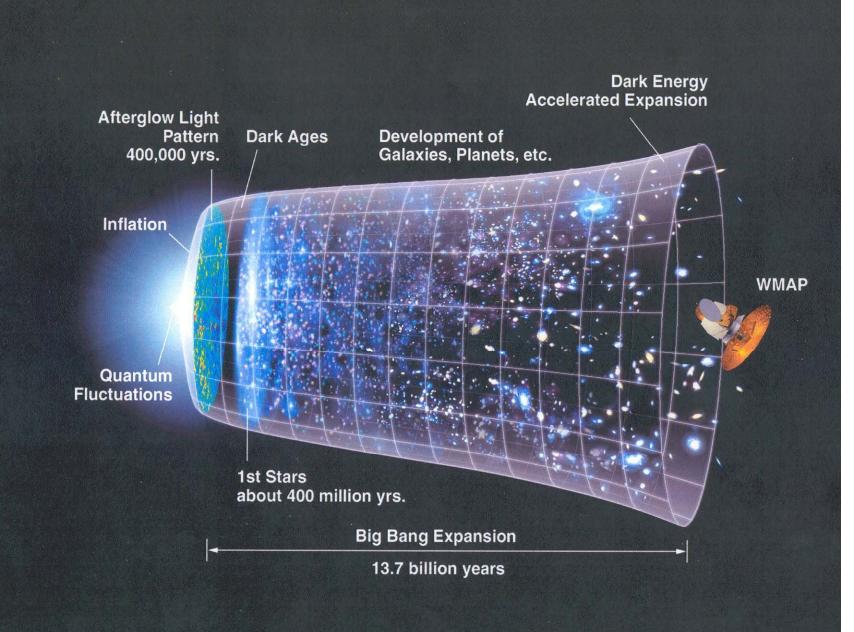


#### The weak force is feeble in the Sun ...

..because 10,000,000K ~ 1 keV << 80 GeV

...this is why the sun has stayed active long enough for us to have evolved and be having this conversation.

- →We exist because m(W) is not zero
- → Mass matters





Thermal Equilibrium:



Temperature (energy) drops >
After 1 usec > one way only:



But at the same time:



then like processes 2 and 3 in the Sun Xuntil all the neutrons have gone

\* making Helium

9

particles so far apart in the expanding universe that they no longer interact

T = 1 usec after BIG BANG

neutrinos are free ("the first fossils in the Universe)

and if they have mass they start clustering together -contribute to formation of galaxies

Billion > per atom

if m(v) > m(proton)/109 ~1eV

they will dominate mass

density of the Universe

or future of universe and its formation

Universe expands - and cools expansion rate



Rate depends on pressure which depends on the temperature in the gas and the number of neutrinos inside the gas volume (density) and this # depends on number of neutrino species

#### T=3 minutes after BIG BANG

75% protons
24%. Helium Nuclei
+ small amount of deuterons
+ free electrons.

Helium abundance\*; Horse of other elements

depends on expansion rate of the Universe which depends on number of neutrino species

Deuterium abundance depends on density of "ordinary matter" in the Universe.

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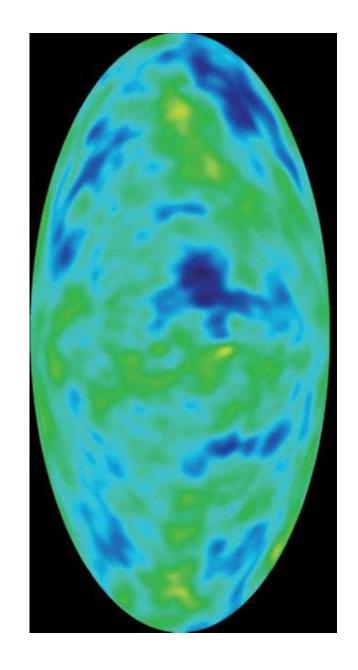
Deuterium abundance

depends on density of "ordinary matter" in the Universe.

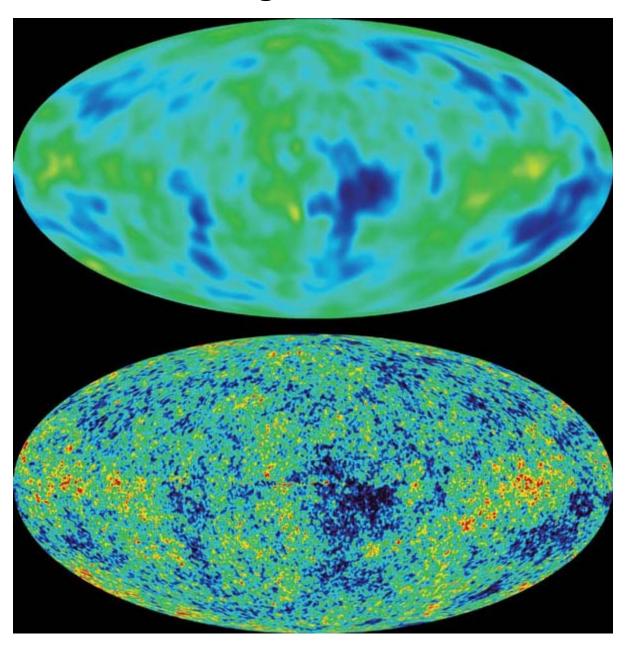
density of ordinary matter <> total in universe

part of DARK MATTER puzzle

Time Passes. Temp drops 300,000 years later ExloeV T< 104K electrons combine with nuclei and make neutral atoms electromagnetic radiation was set free Universe becomes transparent 10 years later Emag 2 stretched: Micronome Band. Black body background 3K (small fluctuations in Hawave rad = hints of proto structures, galaxier in early universe)



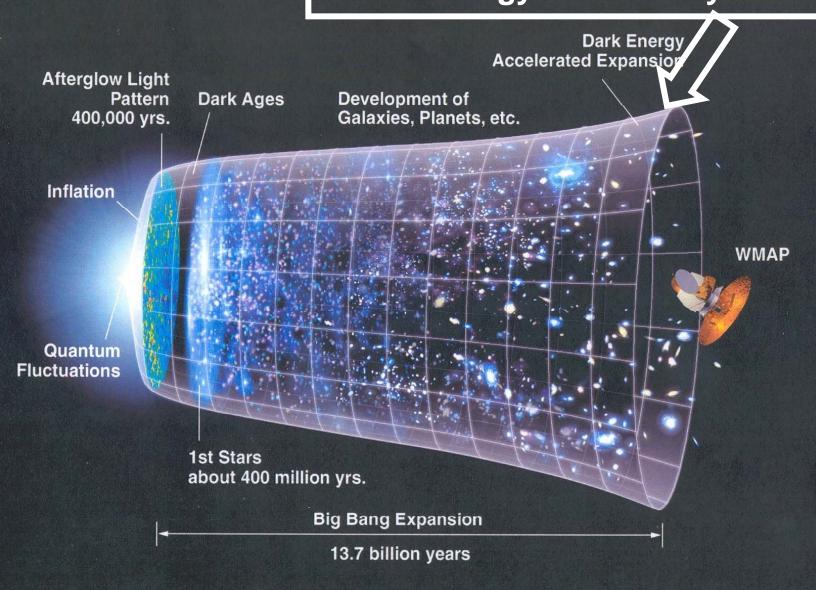
#### 3K microwave bgnd now seen to have structure



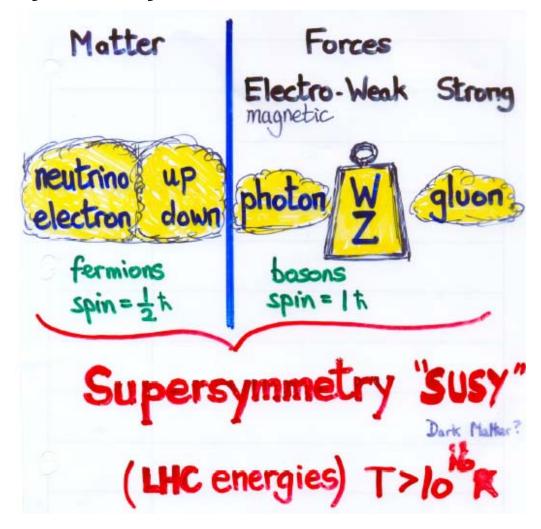
**COBE 2000** 

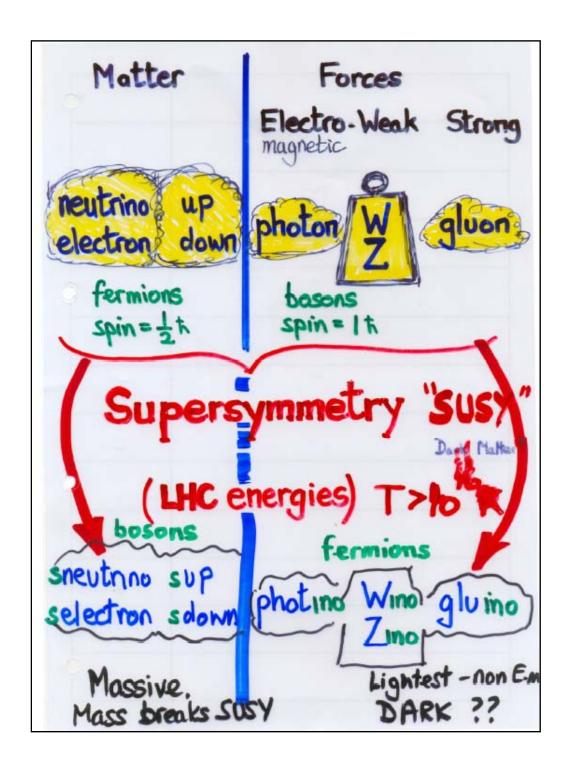
**WMAP 2003** 

### 5byr ago accelerated expansion = Dark Energy. What? Why?.....



#### One further symmetry??





Whole new families to be found

#### Particle Physics @ CERN. Standard Model of PHes + Forces. · Quarks + Leptons. Spin 1/2 fermions · Y WZ gluons Spin 1 gauge basons · Higgs Spin O boson Origins of malter. Structures + pallens at E& ITEV Symmetry revealed at E > O(TeV) Forces (and particles) unified - SUSY: Some current big pazzles. . Dark Malter, Solar v, Massive v? (all the same?) . Why 3 generations What is difference between Mand M? I same? . The Fifth Dimension

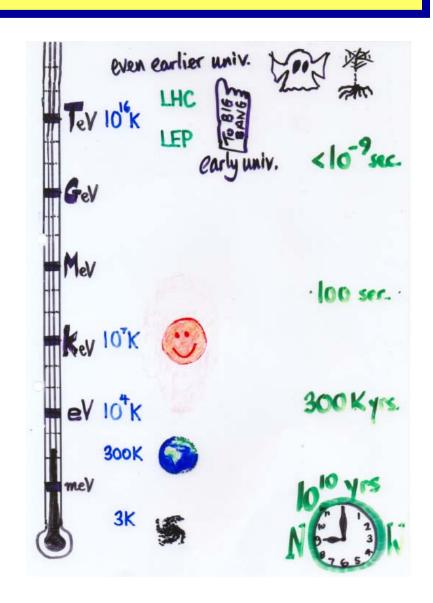
#### Finale: A glimpse of the future

recall from lecture 1.....

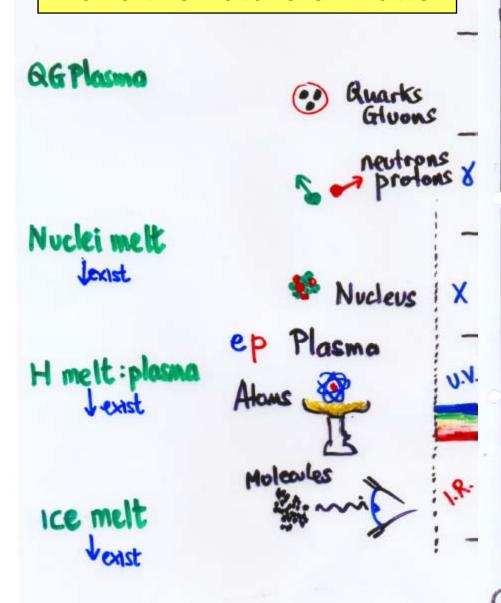
#### **The Universe**

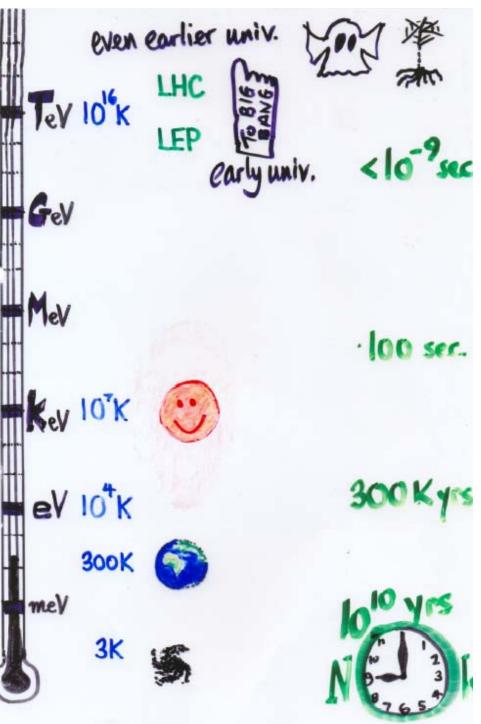
in

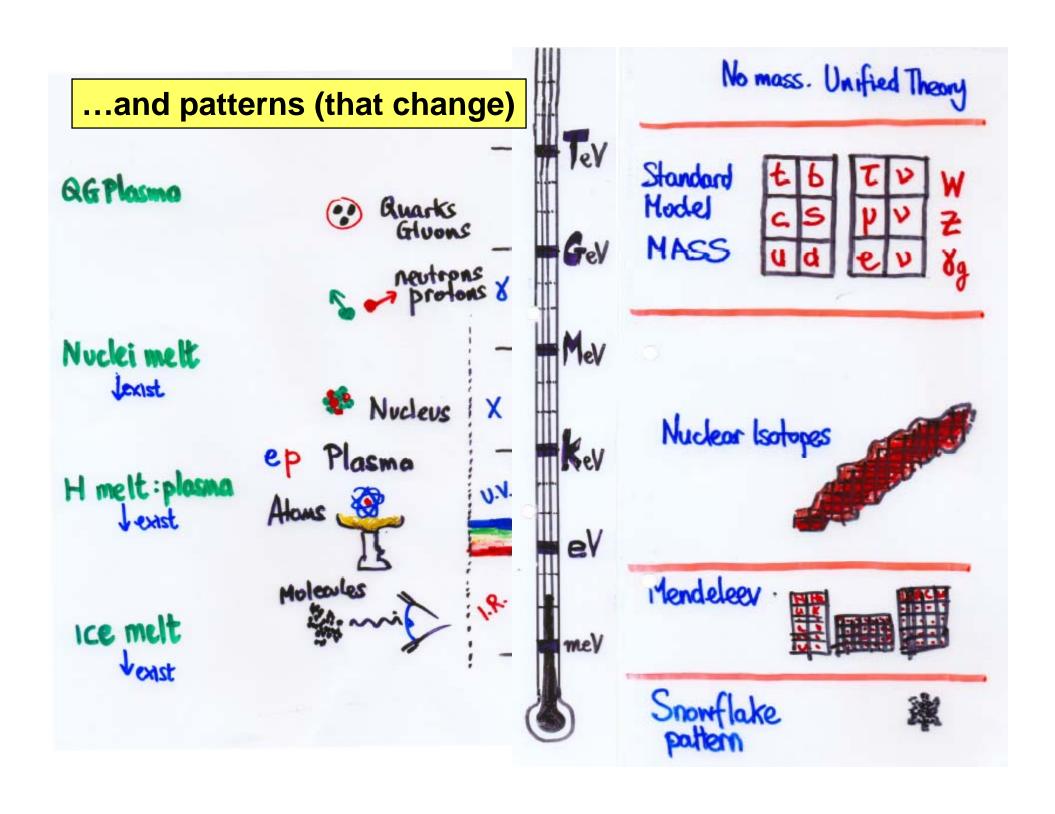
Temperature
Energy and
Time



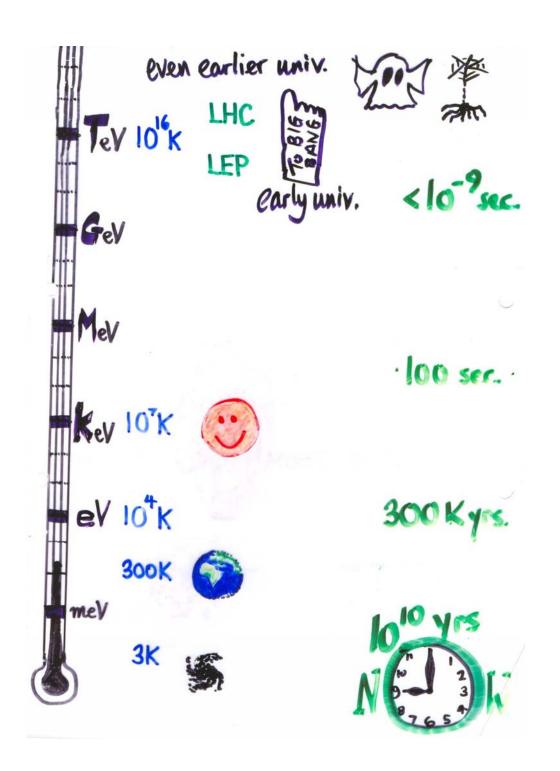
#### ...and the nature of matter







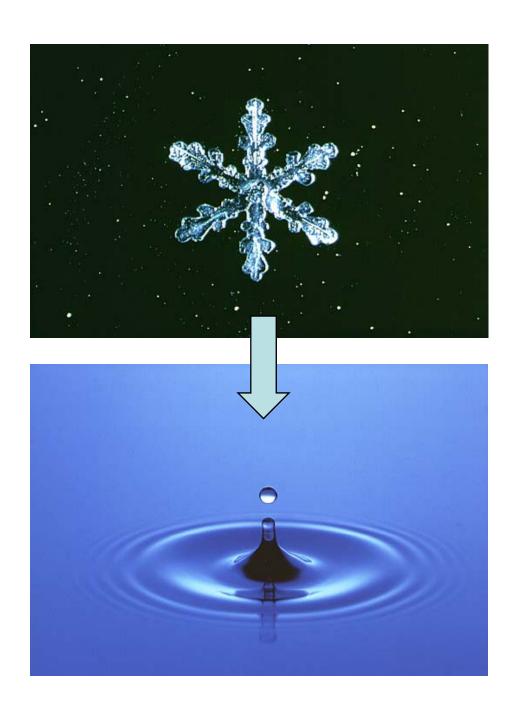
Temperature and symmetry in the universe



## The Idea



(I will tell you when to be cautious about inhaling)

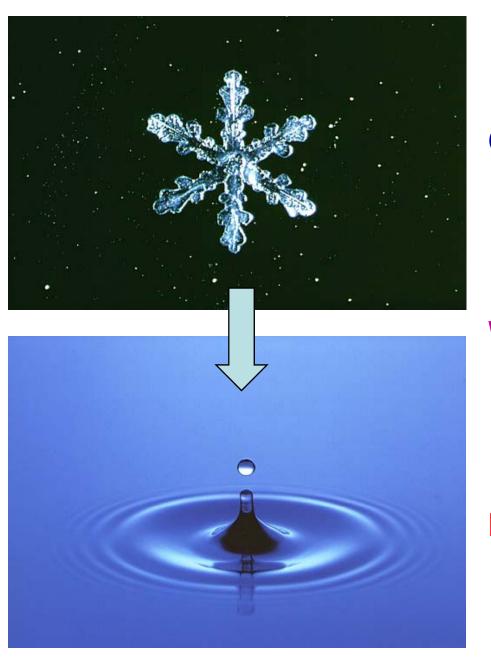


patterns
and structures
when cold
(low energy)

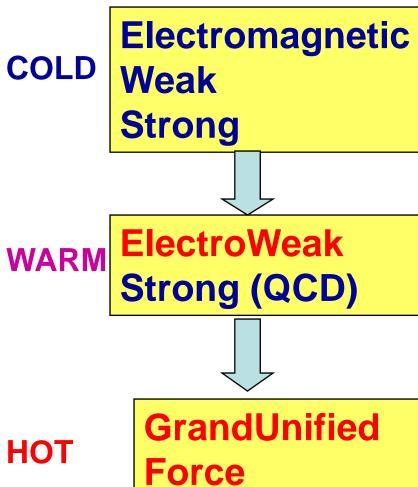
Symmetry
when warm

(high

energy)



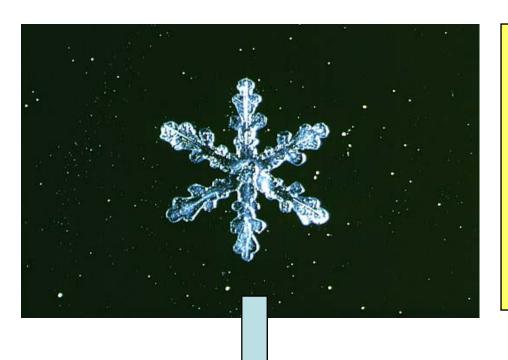
#### **FORCES 1955-2005**





# Standard Model of Quarks Leptons and forces

- = pattern based on Mass
- "cold" ="low" energy
- = below 1 TeV

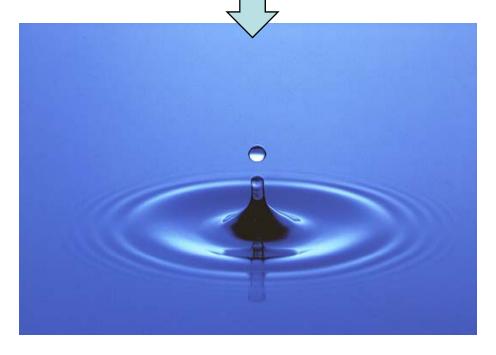


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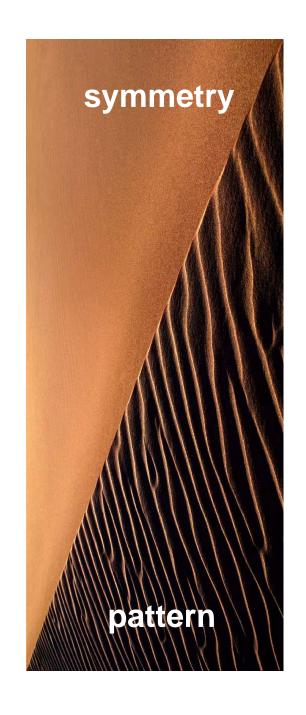
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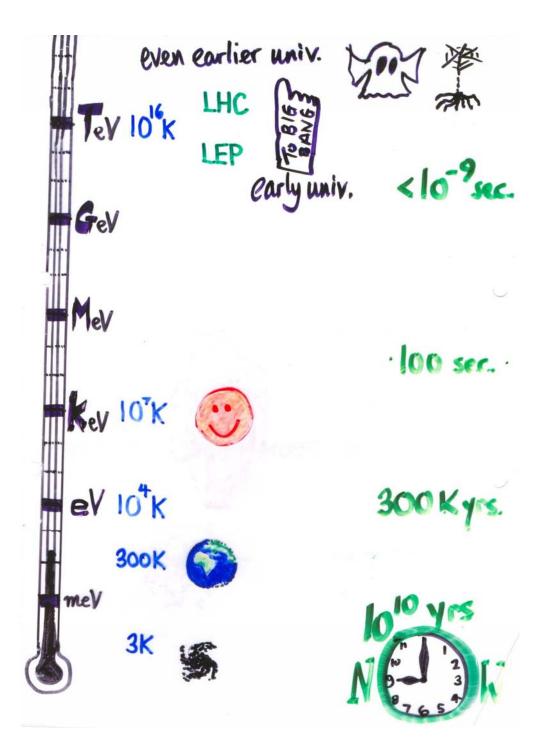
= below 1 TeV



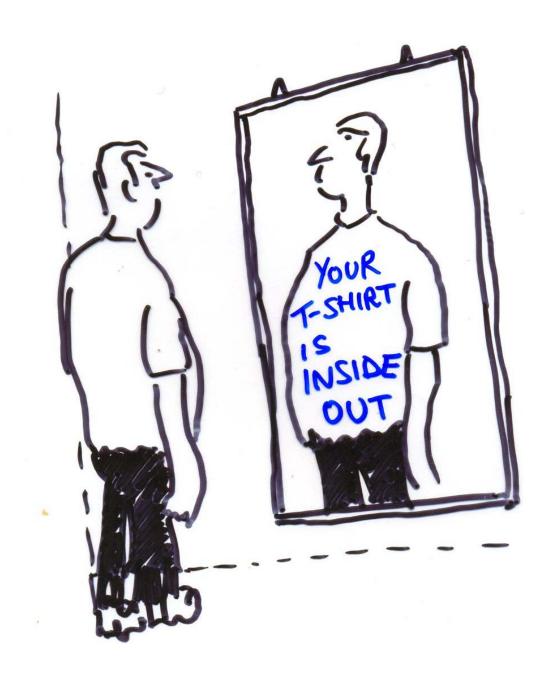
superSymmetry
when "warm"
(= high energy > 1TeV)

Higgs Boson Supersymmetry Nature of Reality



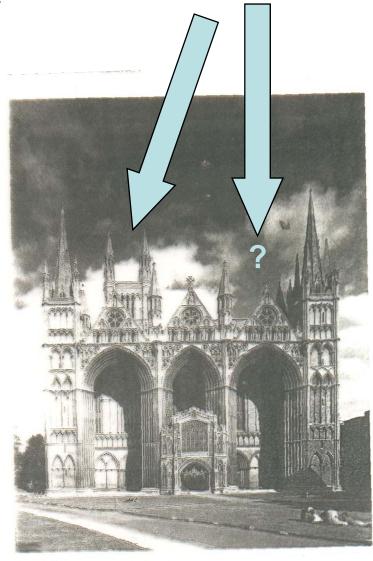


# 5. symmetries can disappear or change



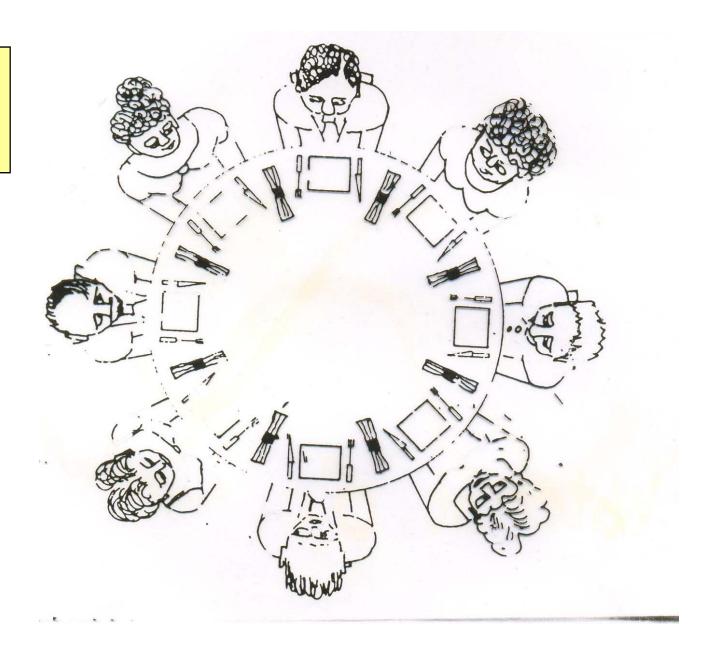
We like symmetry and when its absent we

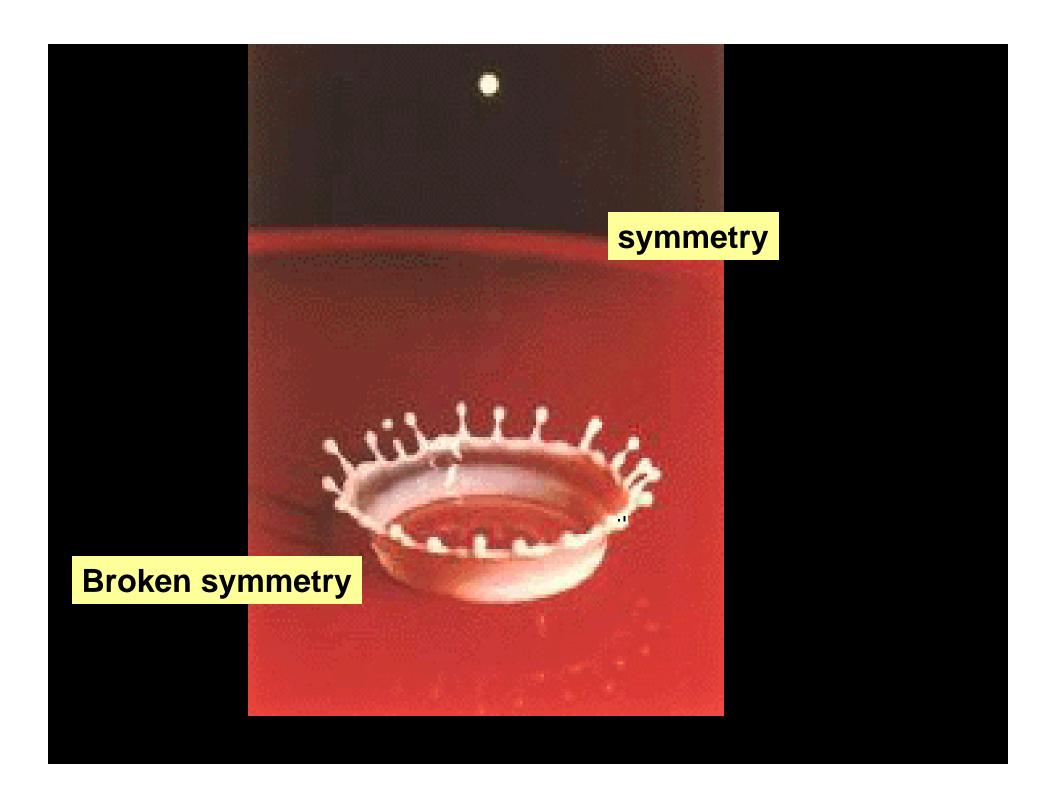
want to know why

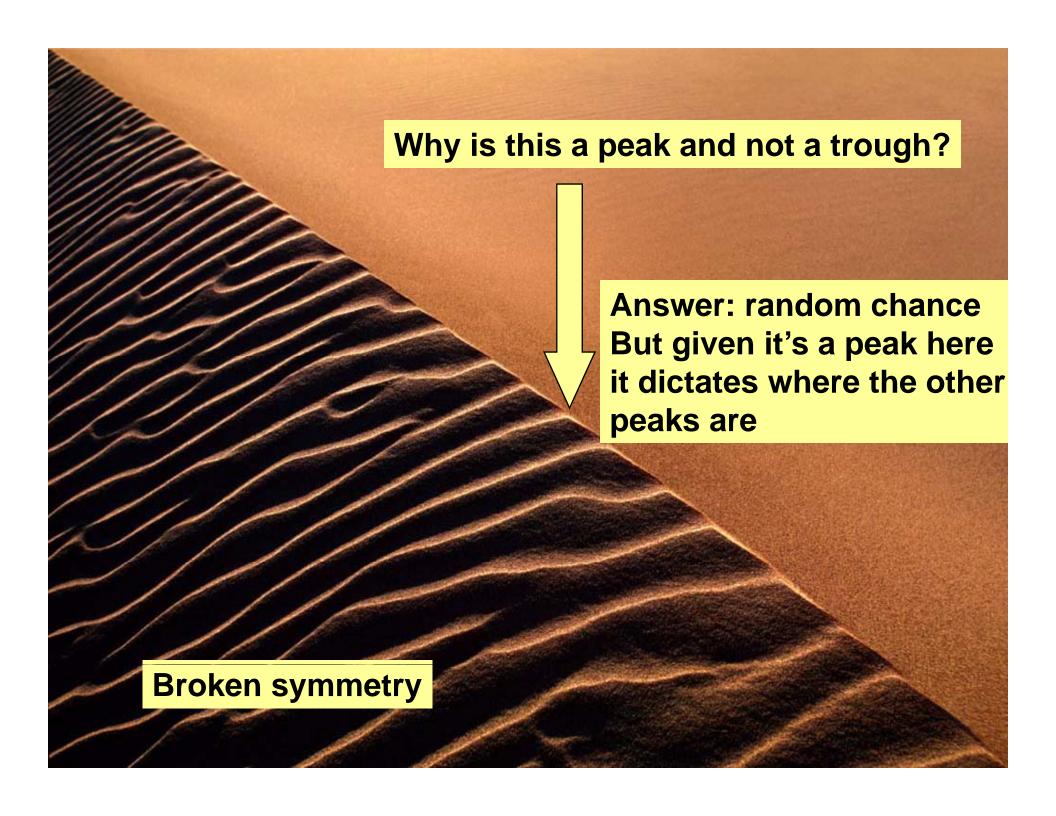


# **Buridan's Ass**

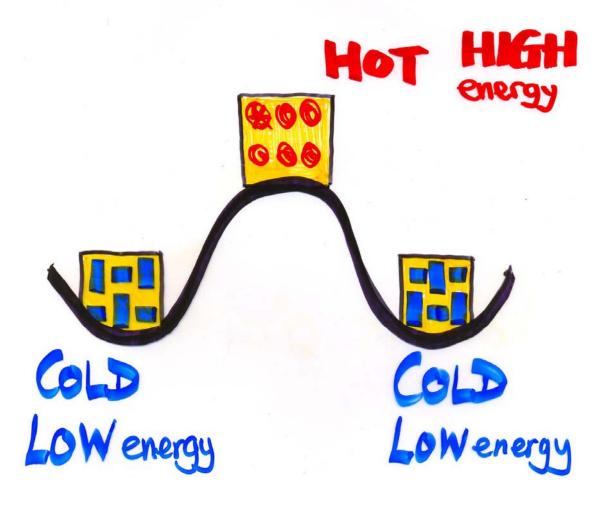
The problem of the symmetric dinner party

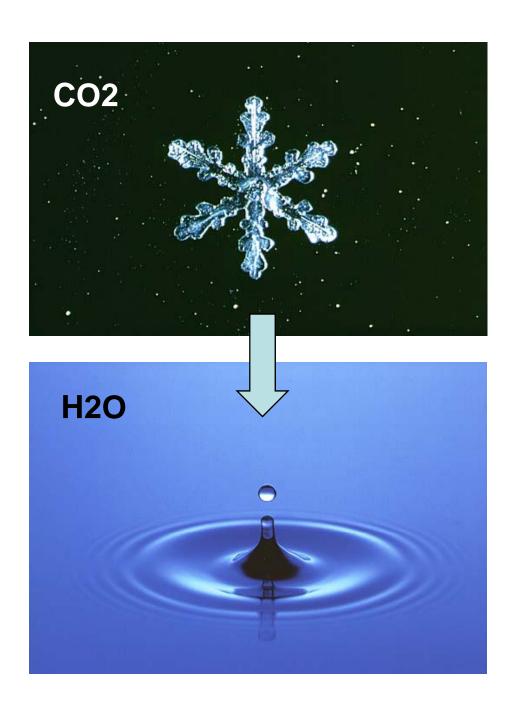






## Magnets





patterns and structures when cold (low energy) **Symmetry** when warm (high

energy)

As the universe cooled after the hot big bang.....
We think that an elegant symmetry......

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... "froze" into structures .... And patterns
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Such as Atoms ..... Mendeleev's periodic table,
And particles .... Quarks, forces and the Standard Model
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... which is a pattern based on MASS

2008: heat up to energies above 1000 GeV = "1 TeV" and discover the origin of MASS (= Higgs?)

