

Dulution of axion dark radiation

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§ Introduction

§ § SM of particle physics

Success of the standard model of particle physics

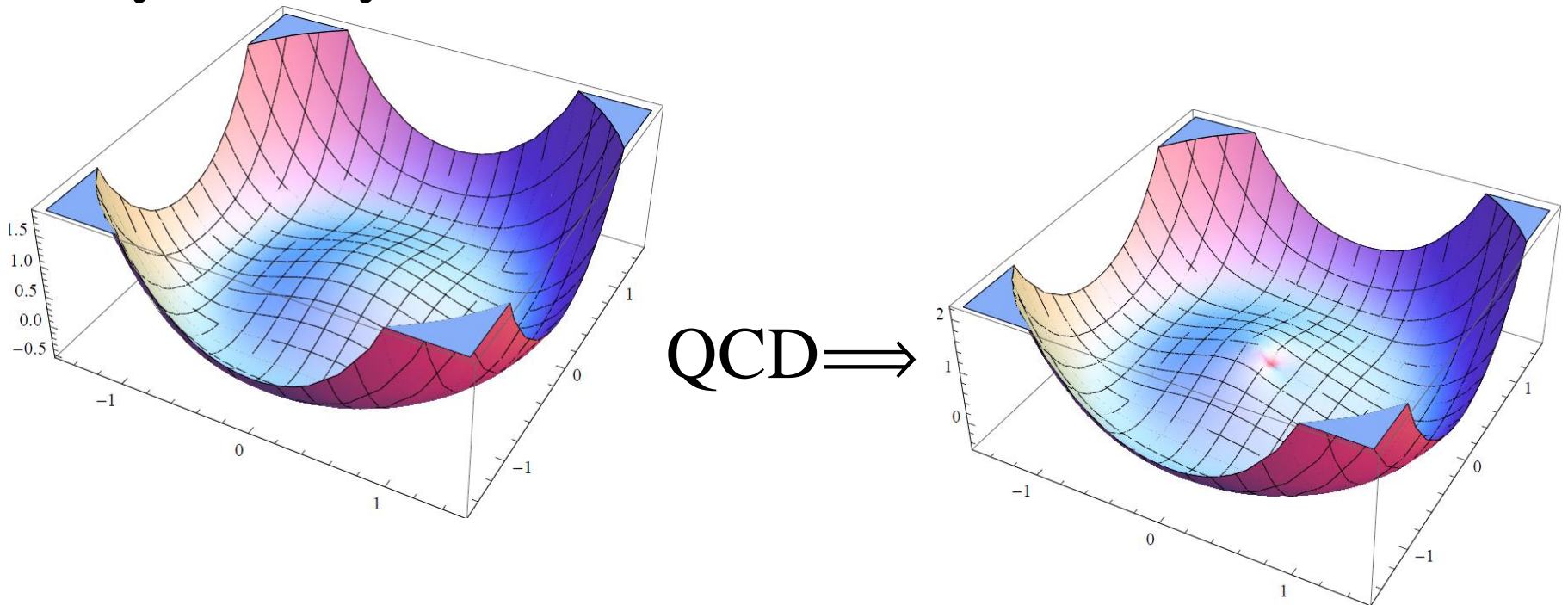
- It well explain almost all phenomena occurred at accelerator experiments.

Problems of the SM

- Nonvanishing neutrino mass
- **Strong CP problem → axion in PQ mechanism**
- Gauge hierarchy problem
- ...

§ § Axion

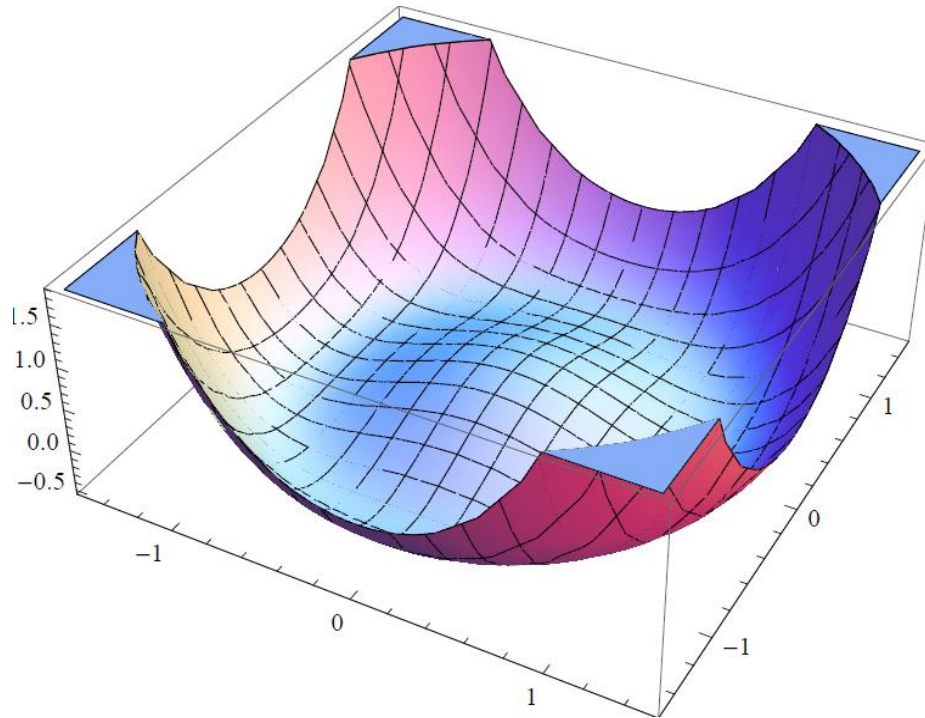
- Axion is a NG boson of broken global $U(1)_{PQ}$ symmetry



- Tiny mass below QH transition.

§ § Axion

- Radial direction of the field also evolves in the early Universe.



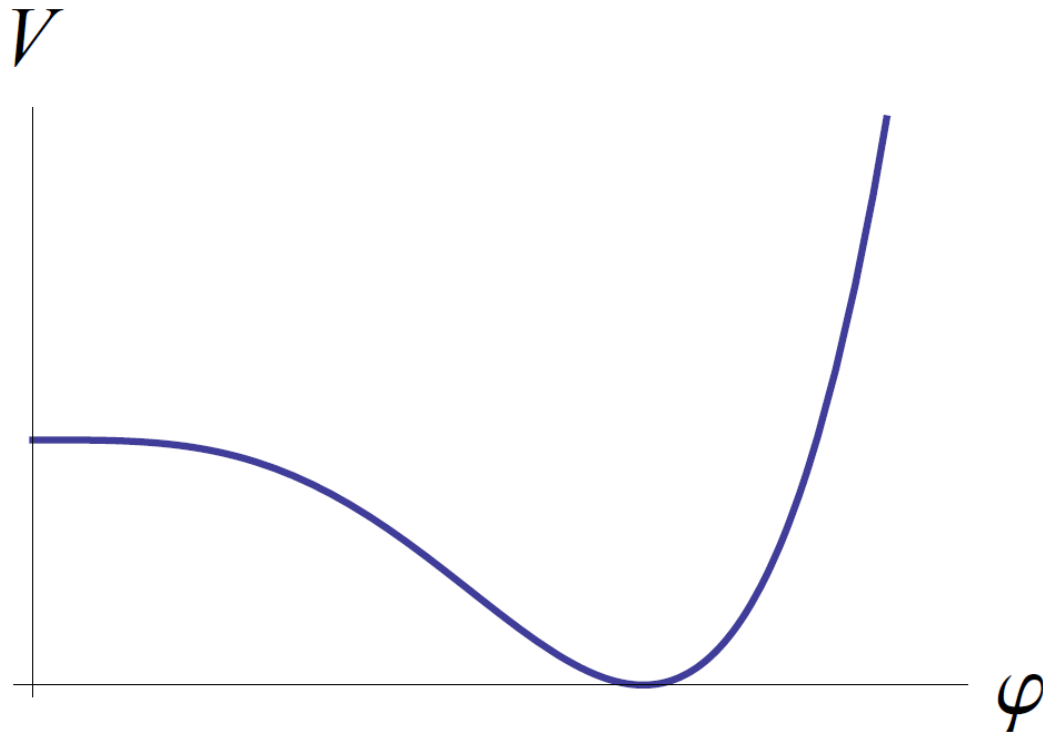
- The radial field (PQ field) would decay into 2axions.

§ Assumed cosmic history

- Primordial inflation
- Reheating by the primordial inflation
- PQ scalar domination
- PQ scalar decay into axions, relativistic axion dominated Universe
- **Thermal inflation**
- **Reheating after thermal inflation**
- **Baryogenesis**, dark matter production, etc
- Big bang nucleosynthesis

§ Thermal inflation

- A promising mechanism to dilute unwanted relics



§ Thermal inflation

- Potential: Higgs for gauged U(1)

$$V(\varphi) = V_0 - m^2|\varphi|^2 + \frac{|\varphi|^{2n}}{\Lambda^{2(n-2)}} + \frac{g_\varphi T^2 |\varphi|^2}{24}$$

- Number of e-fold

$$N_{2n} = N_4 - \frac{1}{4} \ln \frac{n^2}{4(n-1)} + \frac{1}{2}(n-2) \ln \left(\frac{M_P}{v} \right),$$

$$N_4 = -\ln 4\sqrt{3} - \frac{1}{4} \ln \left(\frac{\pi^2}{30} g_* \right) + \frac{1}{2} \ln \frac{\Lambda}{M_P} h,$$

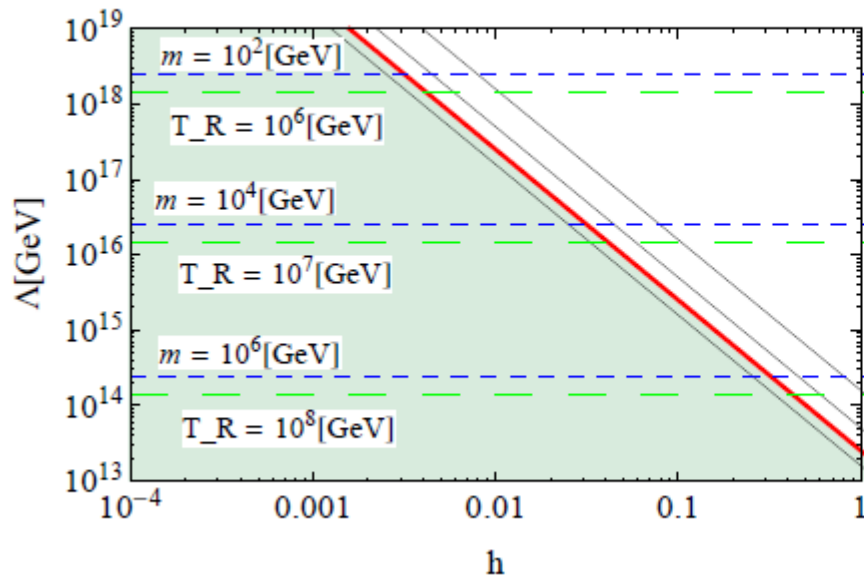
- Results for n=3

TABLE II: Quantities in thermal inflation by the potential (19)

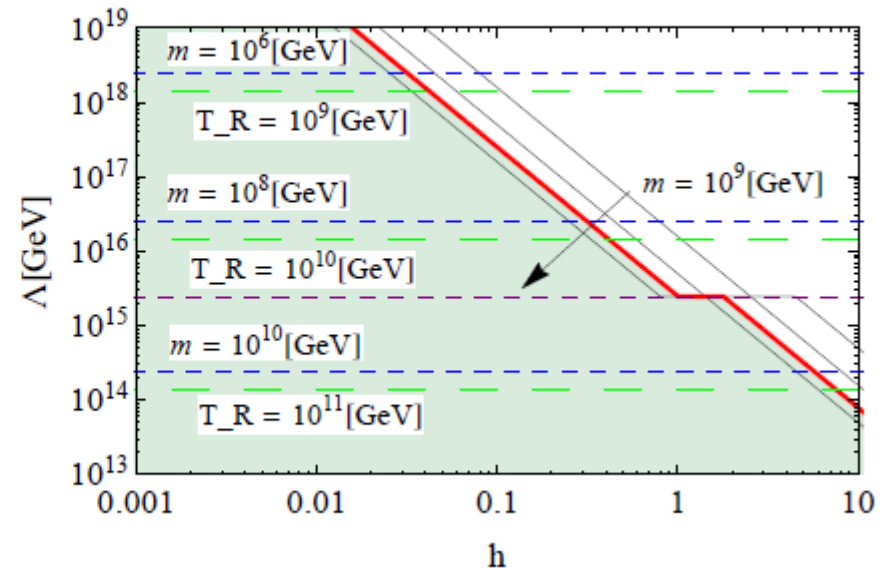
| $\Lambda(\text{GeV})$ | h | $v(\text{GeV})$ | $T_i(\text{GeV})$ | $T_f(\text{GeV})$ | N | ΔN_{eff} | $T_R(\text{GeV})$ |
|-----------------------|-----------------------|-----------------|--------------------|--------------------|------|------------------|-------------------|
| 10^{16} | 8.27×10^{-3} | 10^8 | 2.79×10^3 | 1.03×10^3 | 1.00 | 0.05 | 5.9×10^3 |
| 10^{16} | 8.27×10^{-2} | 10^{10} | 2.79×10^6 | 1.03×10^6 | 1.00 | 0.05 | 5.9×10^6 |
| 10^{16} | 8.27×10^{-1} | 10^{12} | 2.87×10^9 | 1.03×10^9 | 1.00 | 0.05 | 5.9×10^9 |

§ Relic abundance and viable baryogenesis

$n=3, v=10^{10}[\text{GeV}]$



$n=3, v=10^{12}[\text{GeV}]$



Compatible baryogenesis

Low scale thermal leptogenesis

High scale thermal leptogenesis

§ Summary

- PQ scalar might dominate the energy density of the early Universe and decay mostly into relativistic axion.
- Thermal inflation by a Higgs field for a gauged $U(1)$ can dilute them.
- Promising viable baryogenesis: High or low scale thermal leptogenesis