University of Southern Denmark

Compact Stars & Dark Matter

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CP^3 - Origins

Particle Physics & Origin of Mass

MIAPP Munich, 4 February, 2015

Astrophysical Observations

WIMP annihilation and Cooling of Stars

WIMP annihilation as a heating mechanism for
•neutron stars (CK '07, CK Tinyakov '10, Lavallaz Fairbairn '10)
•white dwarfs (Bertone Fairbairn '07, McCullough '10)

WIMP collapse to a Black Hole

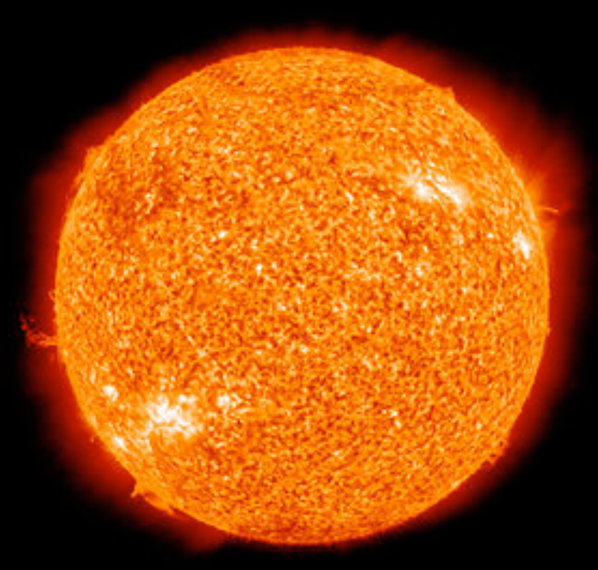
WIMPs can be trapped inside stars and later collapse forming a black hole that destroys the star (Goldman Nussinov '89, CK Tinyakov '10,'11,'13 McDermott Yu Zurek '11, CK'11,'12 Guver Erkoca Reno Sarcevic '12, Fan Yang Chang '12, Bell Melatos Petraki '13, Bramante Fukushima Kumar Stopnitzky '13, Bramante Linden '14, Autzen CK '14)

New effects

WIMPs can slow down the rotation of a pulsar (Perez-Garcia,CK '14)

WIMP capture in Stars

Condition: The energy loss in the collision should be larger than the asymptotic kinetic energy of the WIMP far out of the star.

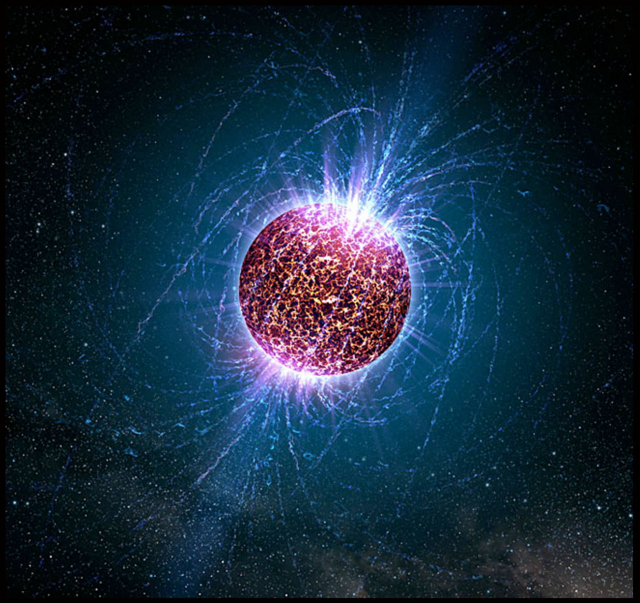


Example: Sun

WIMP mean free path inside the sun $\xi \approx \frac{1}{n\sigma}, \quad n \approx \frac{M_{solar}}{(4/3)\pi R_{solar}^3 m_n} \approx 8 \cdot 10^{23} \quad particles/cm^3$

Even if current limit of CDMS $\sigma < 10^{-41} cm^2, \quad \xi \approx 10^{17} cm, \quad \frac{R_{solar}}{\xi} \approx 10^{-6}$

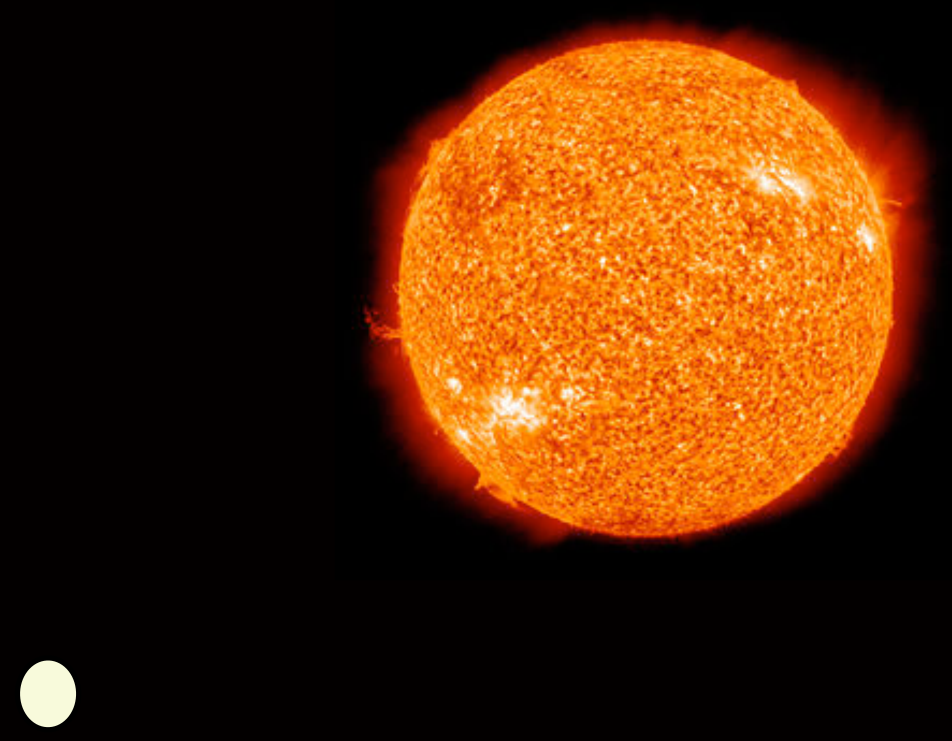
Only one out of a million WIMPs scatters!

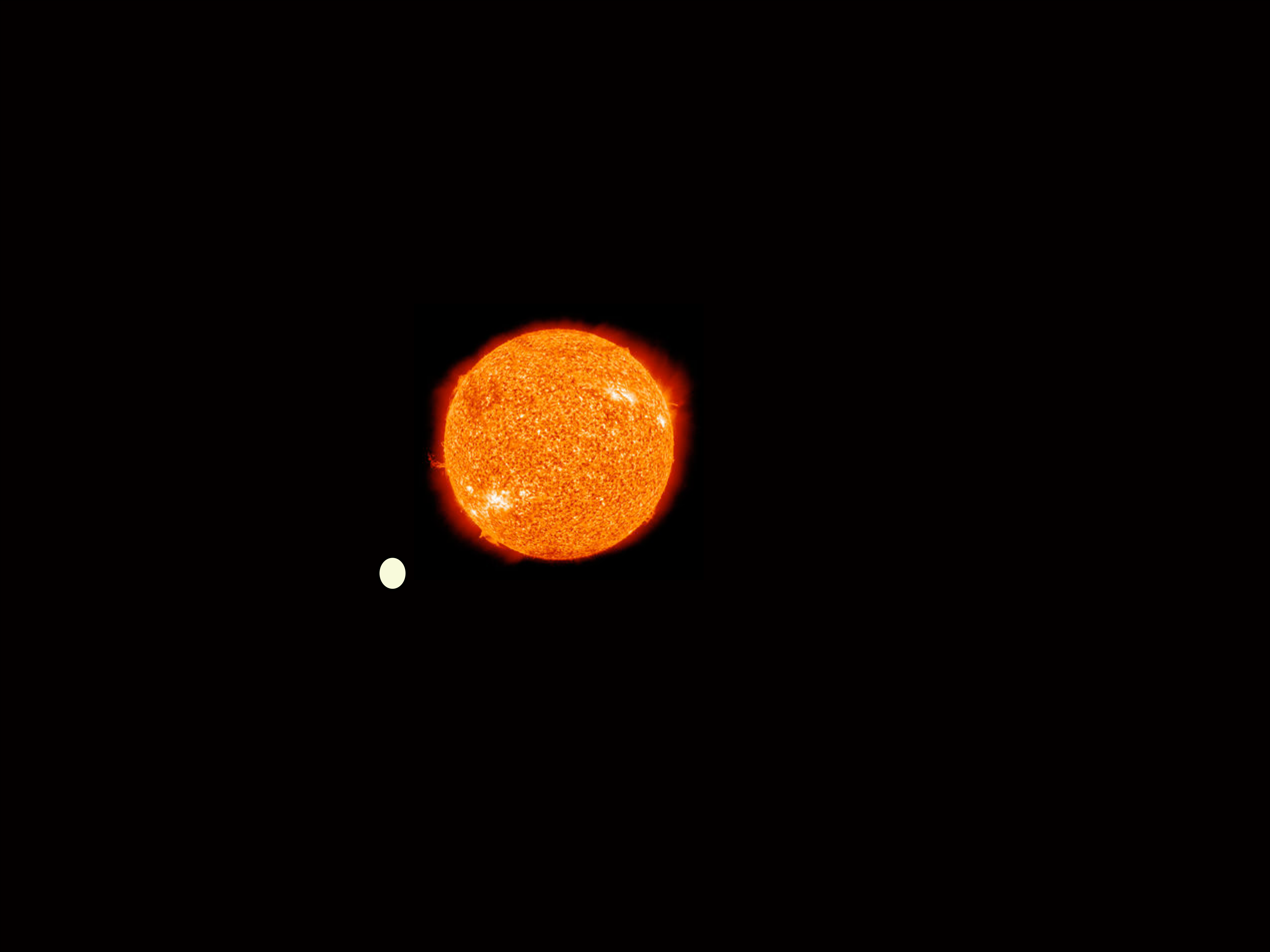


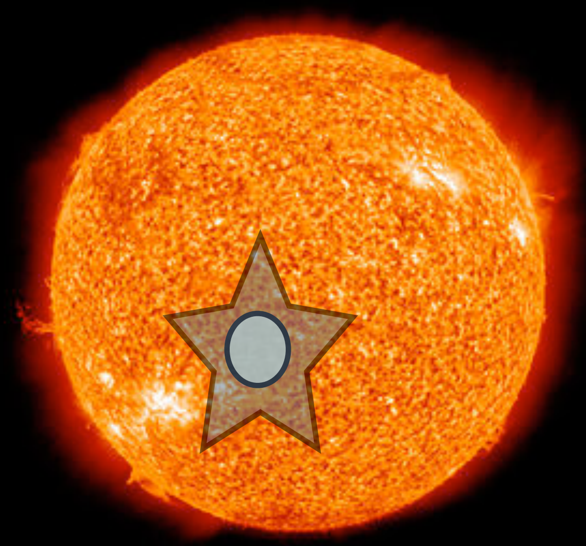
For a typical neutron star $M_{NS} \approx 1.4 M_{solar}, \quad R \approx 10 km$

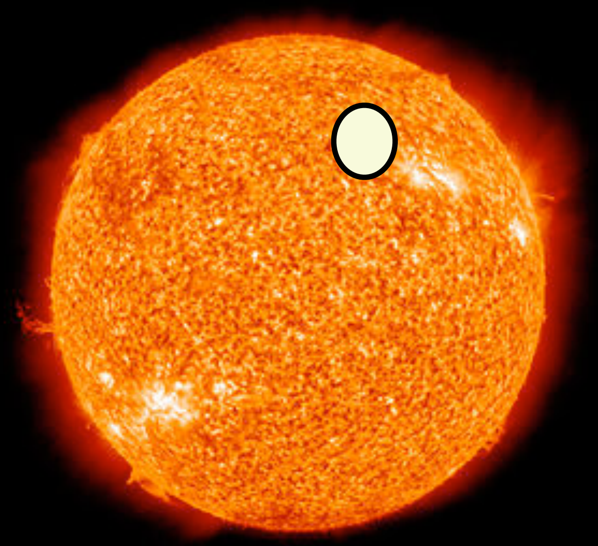
$$\sigma > \sigma_{critical} \approx 5 \cdot 10^{-46} cm^2$$ CK'07

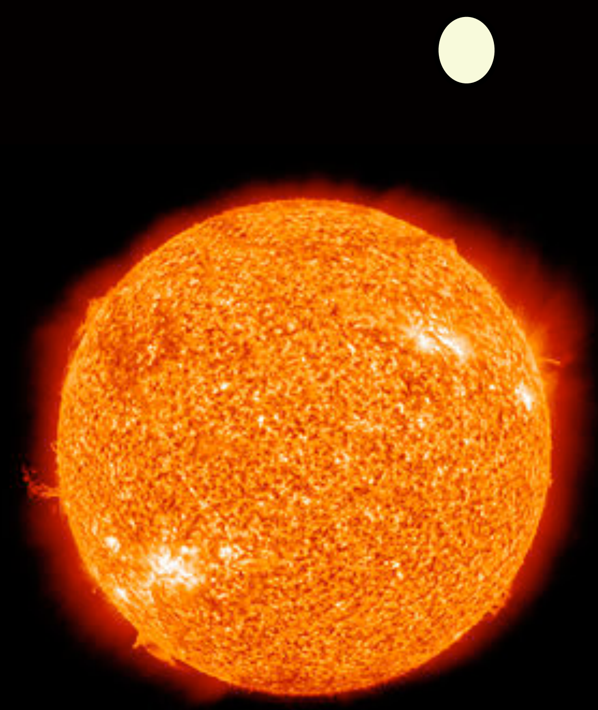
For cross section larger than the critical one, every WIMP passing through the neutron star will be on average interact inside the star.

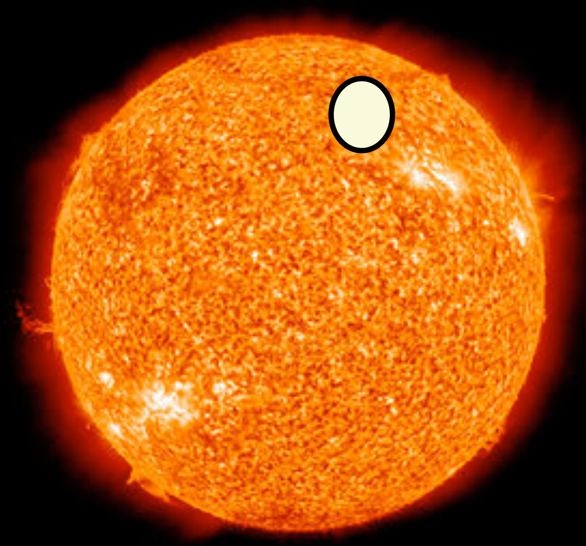


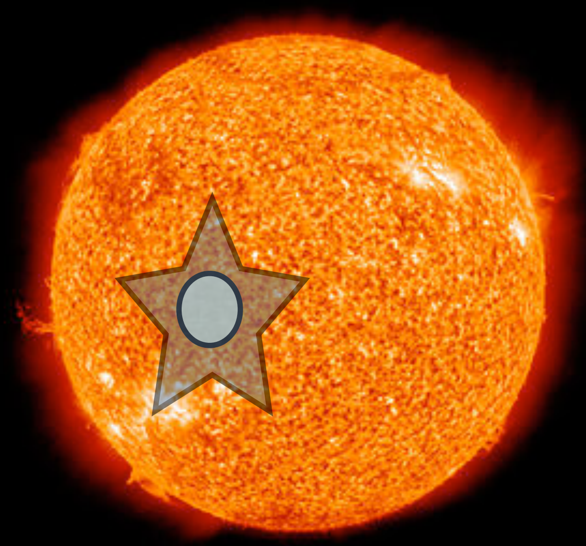


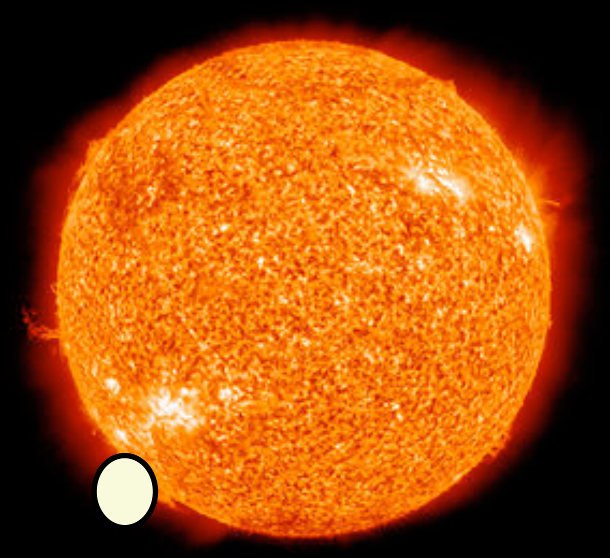


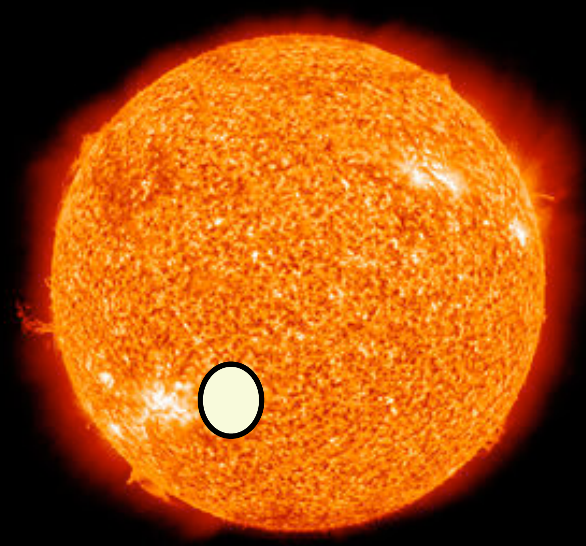


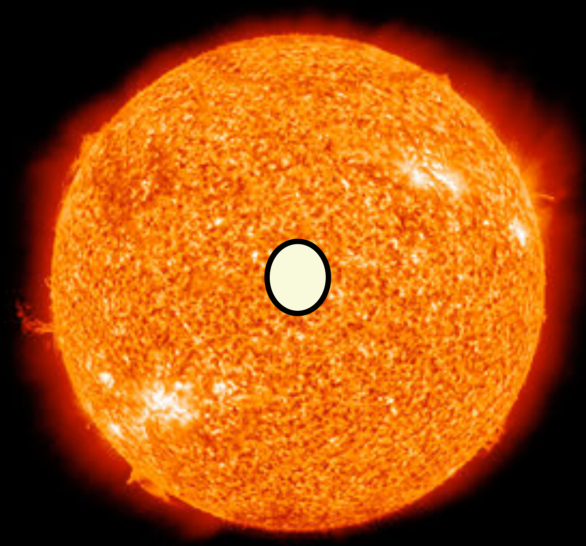


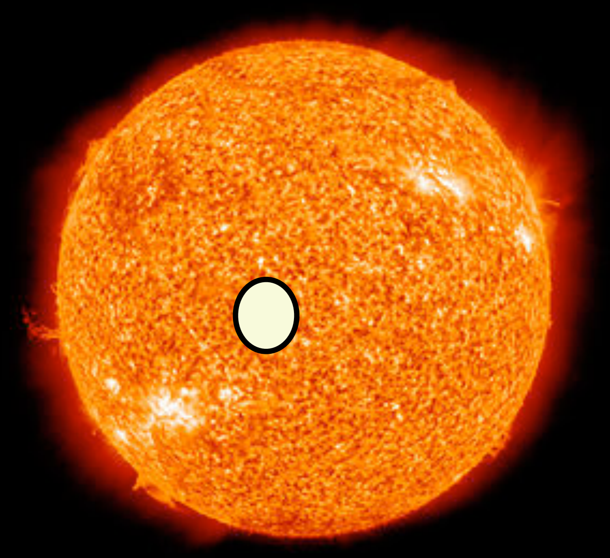


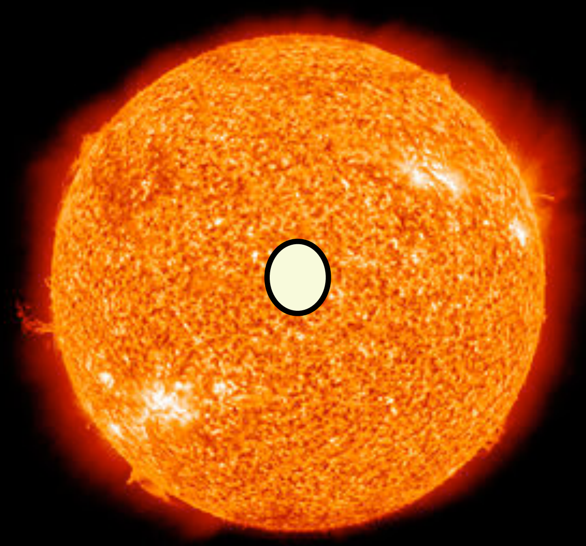


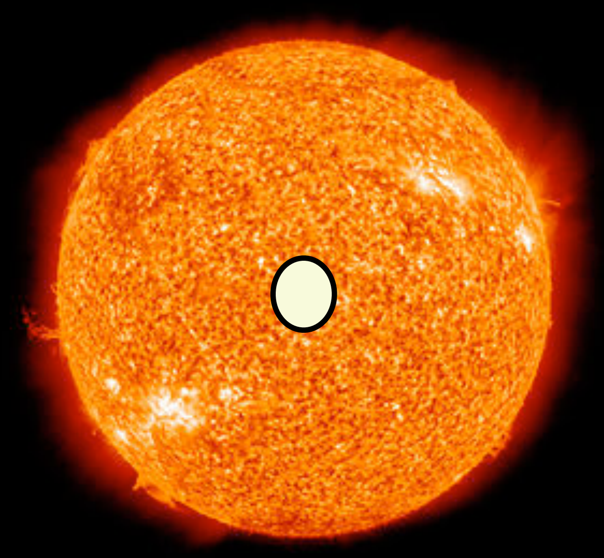


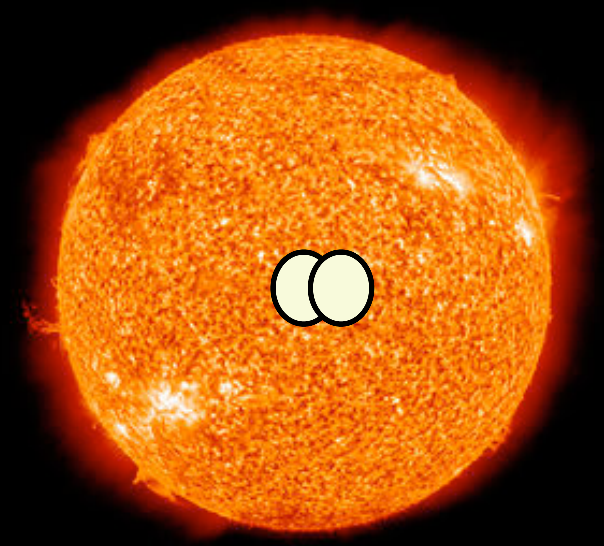


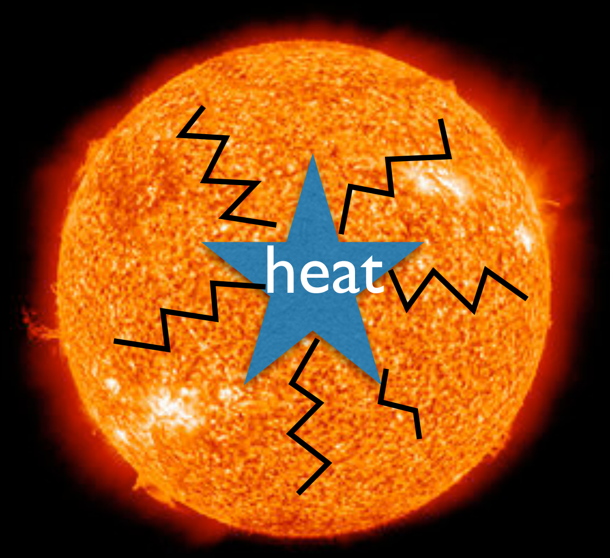











heat

WIMP capture in Stars

$$F = \frac{8}{3}\pi^4 \frac{\rho_{\rm dm}}{m}\left(\frac{3}{2\pi v^2}\right)^{3/2} \frac{GMR}{1-\frac{2GM}{R}} v^2 \left(1 - e^{-3E_0/v^2}\right) f$$

Press Spergel ’85, Gould ’86, Nussinov Goldman ’89, CK’07, CK Tinyakov ’10

higher local DM density gives higher accretion

smaller velocities enhance capture

$f=1$ if $\sigma>\sigma_{crit}$

$f=0.45\sigma/\sigma_{crit}$ if $\sigma<\sigma_{crit}$

$$f \simeq \frac{\sigma_\chi}{\sigma_{\rm crit}} \left\langle \int \frac{\rho}{M/R^3}\frac{dl}{R} \right\rangle$$

For typical NS

$$F = 1.25\times 10^{24}{\rm s}^{-1}\left(\frac{\rho_{\rm dm}}{\rm GeV/cm^3}\right)\left(\frac{100{\rm GeV}}{m}\right) f$$

Thermalization

$$t_{\rm th} = 0.2{\rm yr}\left(\frac{m}{\rm TeV}\right)^2\left(\frac{\sigma}{10^{-43}{\rm cm}^2}\right)^{-1}\left(\frac{T}{10^5{\rm K}}\right)^{-1}$$

Goldman
Nussinov'89,
CK Tinyakov '10
Bertoni Nelson
Reddy '13

$$r_{\rm th} = \left(\frac{9T}{8\pi G\rho_c m}\right)^{1/2} \simeq 22{\rm cm}\left(\frac{T}{10^5{\rm K}}\right)^{1/2}\left(\frac{100{\rm GeV}}{m}\right)^{1/2}$$

Evaporation

$$F = n_s\left(\frac{T}{2\pi m}\right)^{1/2}\left(1+\frac{GMm}{RT}\right)\exp\left(-\frac{GMm}{RT}\right)$$

Krauss Srednicki
Wilczek '86

for WIMPs with mass larger than ~2 keV evaporation can be ignored

WIMP Annihilation in Neutron Stars

$$C_A = \langle \sigma_A v \rangle / V$$

$$\tau = 1/\sqrt{F C_A}$$

$$\tau = 3.4 \times 10^{-5}\mathrm{yr} \left(\frac{100}{m}\right)^{1/4} \left(\frac{\mathrm{GeV/cm^3}}{\rho_{\mathrm{dm}}}\right)^{1/2} \left(\frac{10^{-36}\mathrm{cm}^2}{\langle \sigma v \rangle}\right)^{1/2} \left(\frac{T}{10^5\mathrm{K}}\right)^{3/4} f^{-1/2}$$

Energy Release $$W(t) = Fm\,\mathrm{Tanh}^2 \frac{t+c}{\tau}$$

we have to compare with other heating/cooling mechanisms

Basics of Neutron Star Cooling

Urca process

Direct Urca

$$n \rightarrow p + e + \bar{\nu}_e$$
$$p + e \rightarrow n + \nu_e$$

However for nuclear matter triangle inequalities are not satisfied

For quark matter it holds!

Emissivity: $\propto T^6$

Modified Urca presence of bystander

$$n + n \rightarrow n + p + e + \bar{\nu}_e$$
$$p + e + n \rightarrow n + n + \nu_e$$

Emissivity: $\propto T^8$

Photon Emission Emissivity: $\propto T^4$

$$T_{\mathrm{surface}} = (0.87 \times 10^6\ \mathrm{K}) \left(\frac{g_s}{10^{14}\mathrm{cm/s^2}}\right)^{1/4} \left(\frac{T}{10^8\mathrm{K}}\right)^{0.55}$$

... more cooling mechanisms

- Nucleon Pair Bremsstrahlung

$$n + n \rightarrow n + n + \nu + \bar{\nu}$$

$$n + p \rightarrow n + p + \nu + \bar{\nu}$$

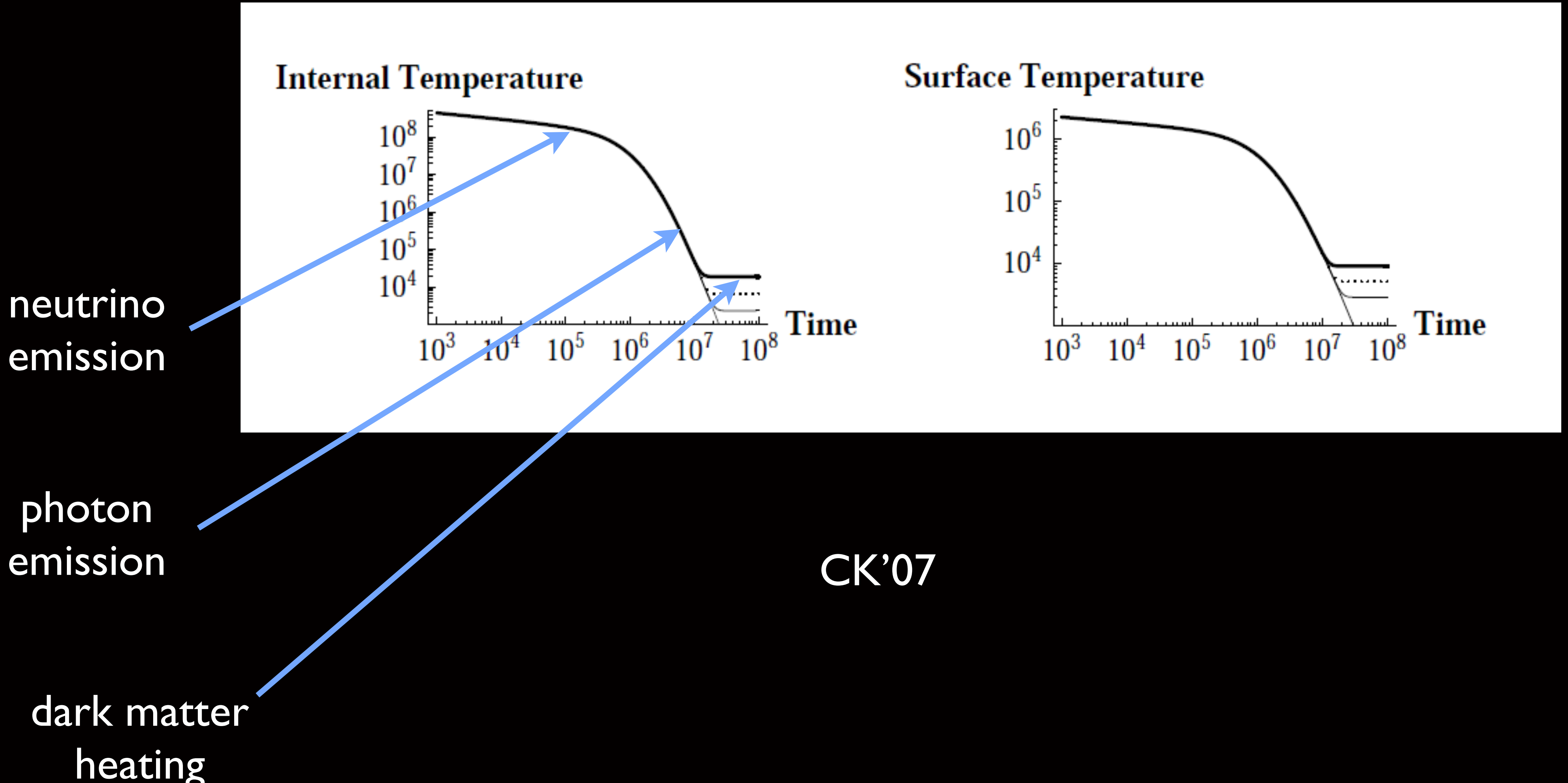
- Neutrino Pair Bremsstrahlung

$$e + (A, Z) \rightarrow e + (A, Z) + \nu + \bar{\nu}$$

- Pionic Reactions
- Superfluidity
- Color Superconductivity

Cooling of Neutron Stars

$$\frac{dT}{dt} = \frac{-L_\nu - L_\gamma + L_{\rm dm}}{V c_V} = \frac{V(-\epsilon_\nu - \epsilon_\gamma + \epsilon_{\rm dm})}{V c_V} = \frac{-\epsilon_\nu - \epsilon_\gamma + \epsilon_{\rm dm}}{c_V}$$



CK'07

Cooling of Neutron Stars

Galactic Center

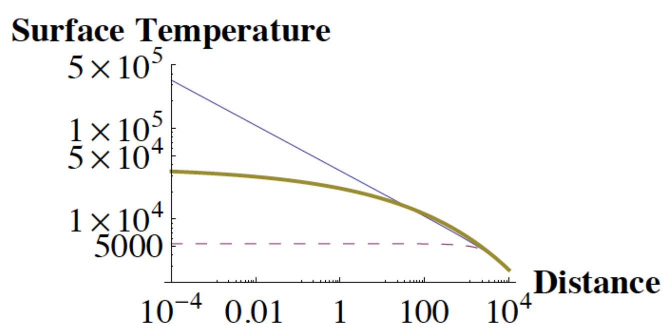


FIG. 3: The surface temperature of a typical old neutron star in units of K as a function of the distance of the star from the galactic center in pc, with the dark matter annihilation taken into account. The three curves correspond to three different dark matter profiles: NFW (thin solid line), Einasto (thick solid line), and Burkert (dashed line).

Globular Cluster

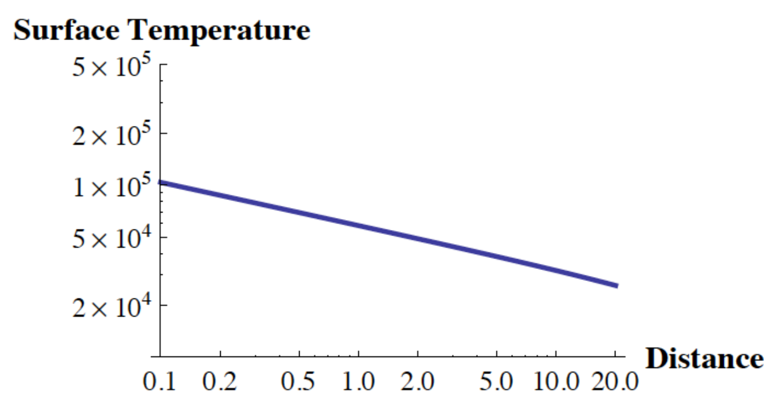


FIG. 5: The surface temperature of a typical old neutron star in units of K as a function of the distance in pc for a NFW profile of the globular cluster $M4$.

$$\rho_{\rm NFW} = \frac{\rho_s}{\frac{r}{r_s}(1+\frac{r}{r_s})^2}$$

$$\rho_{\rm Ein} = \rho_s \exp\left[-\frac{2}{\alpha}\left[\left(\frac{r}{r_s}\right)^{\alpha} - 1\right]\right]$$

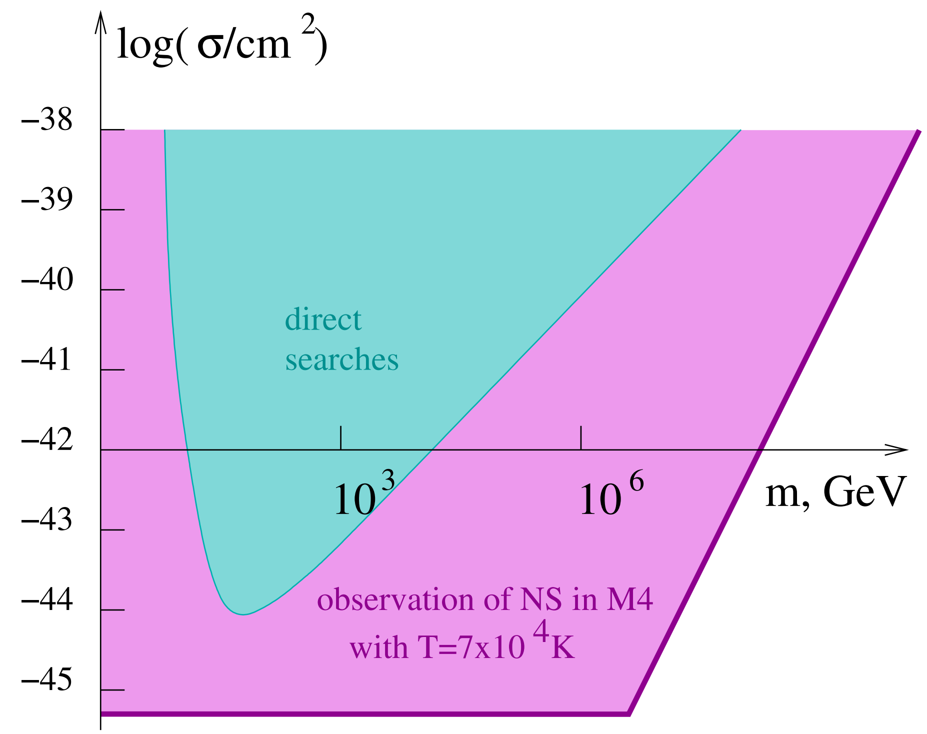
$$\rho_{\rm Bur} = \frac{\rho_s}{\left(1+\frac{r}{r_s}\right)\left[1+\left(\frac{r}{r_s}\right)^2\right]}$$

CK, Tinyakov '10
Fairbairn Lavallaz'10

Nearby old neutron stars

J0437-4715 temperature ~10^5 K
J2124-3358 temperature ~10^5 K
130-140 pc away

Cooling of Neutron Stars

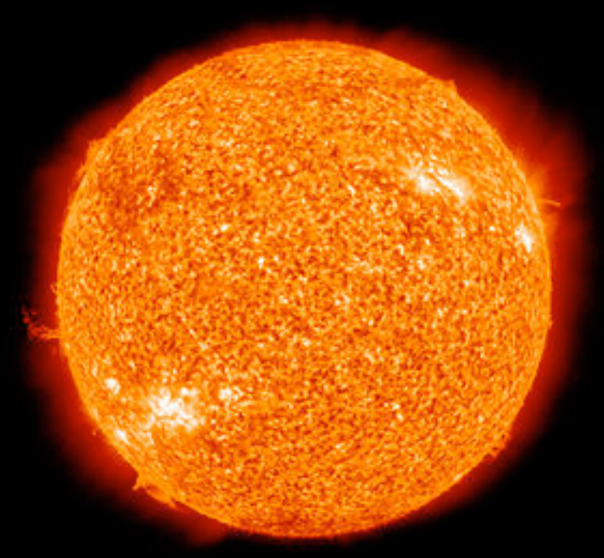


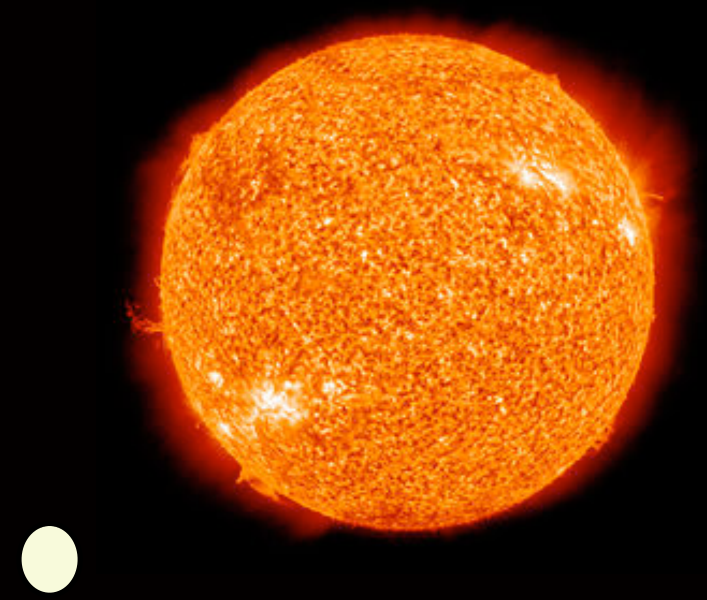
Old neutron stars in Globular Clusters

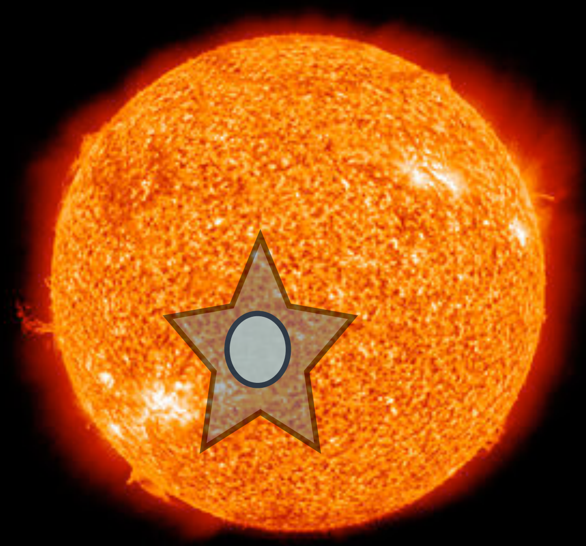
X7 in 47 Tuc
1620-26 in M4

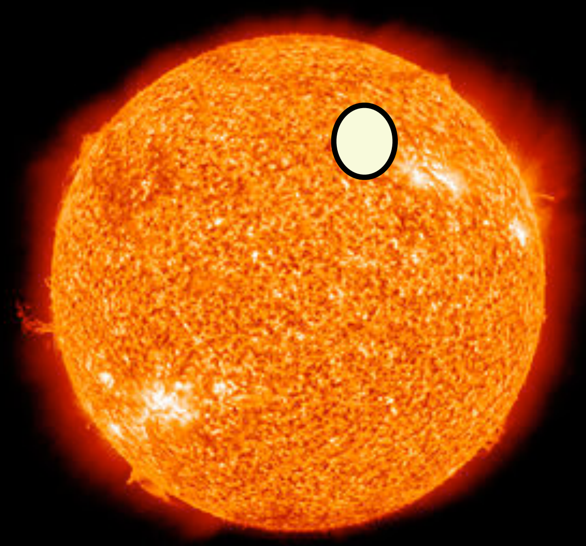
both have temperatures roughly 10^6 K

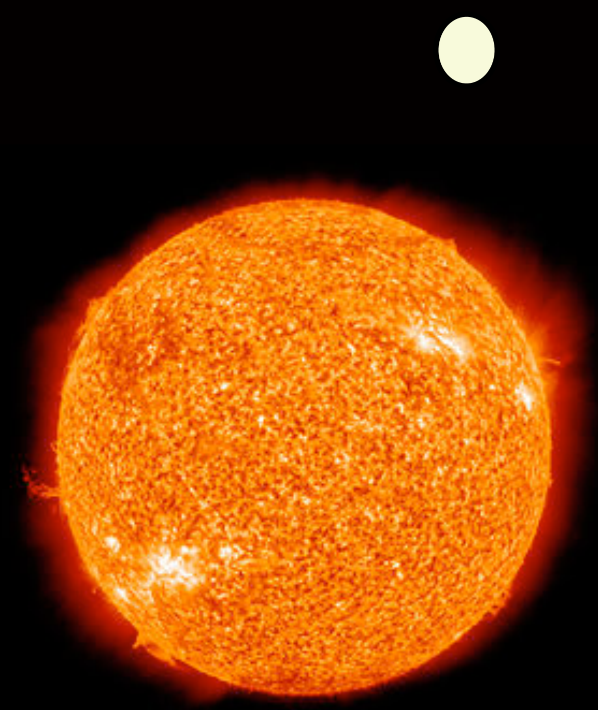
Another scenario for asymmetric DM

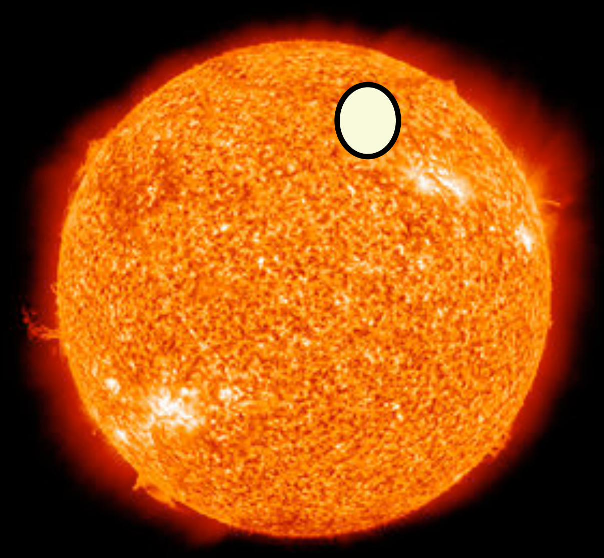


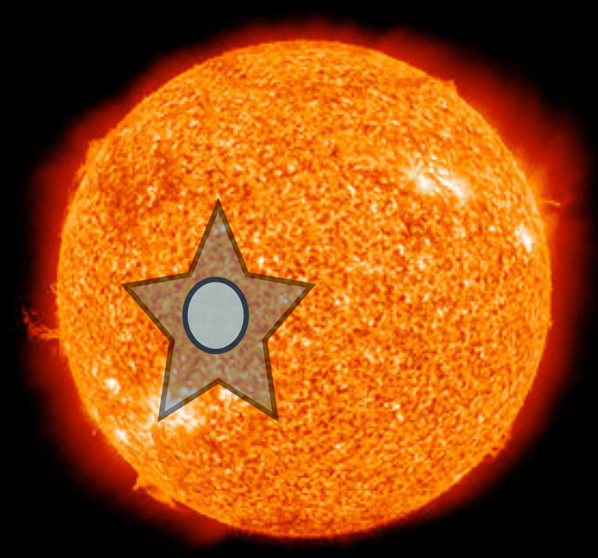


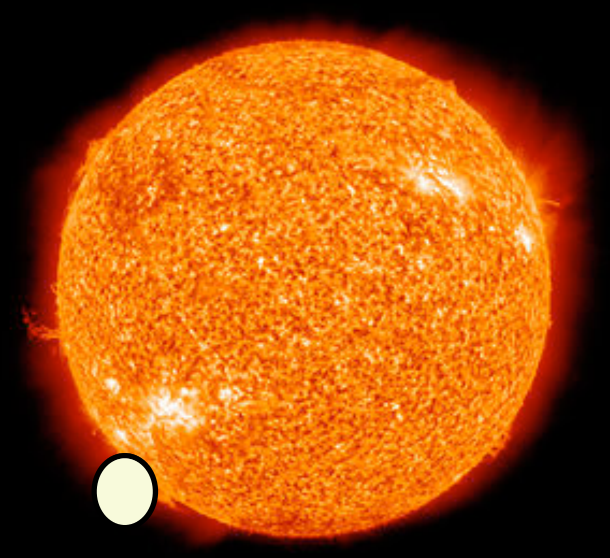


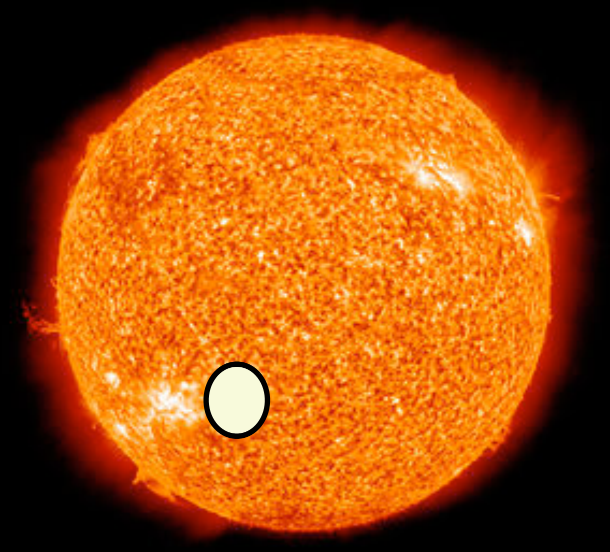


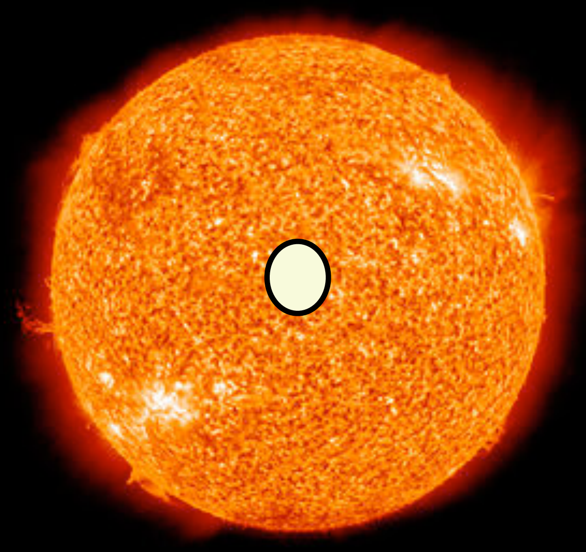


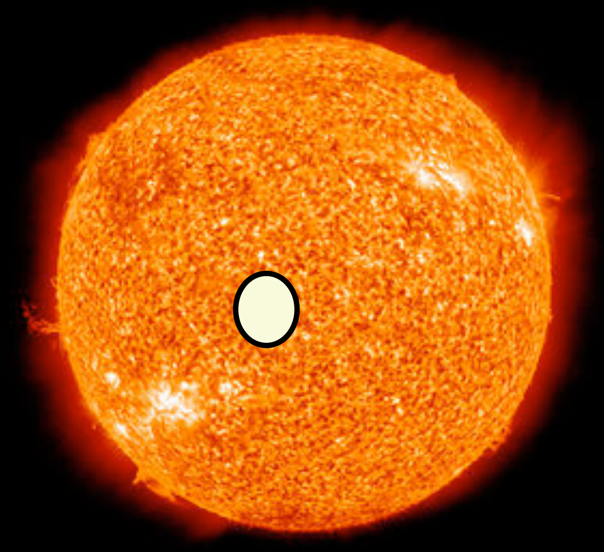


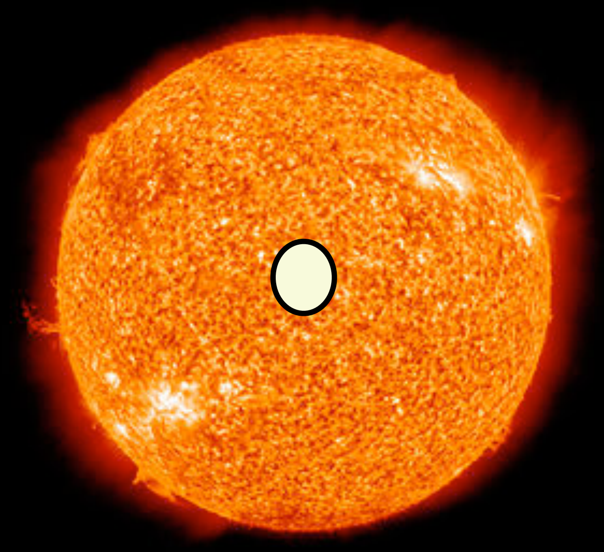


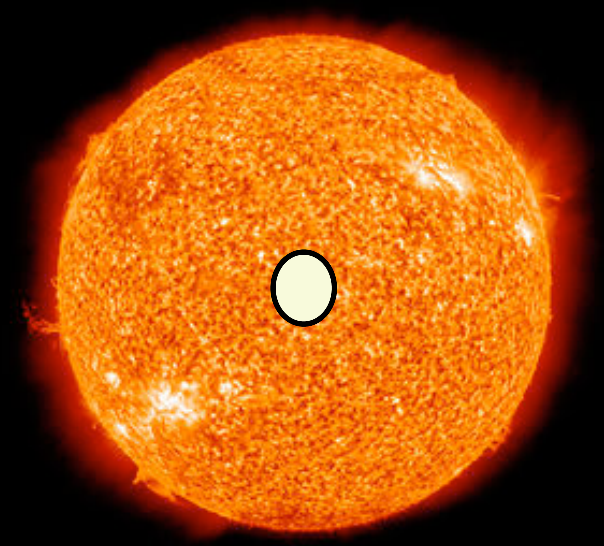


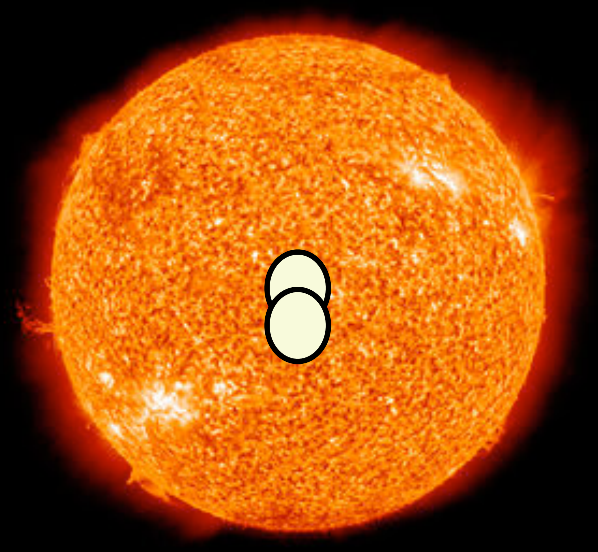


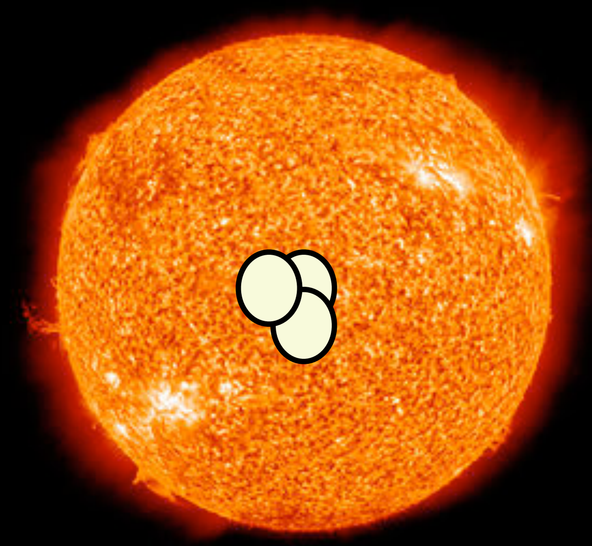


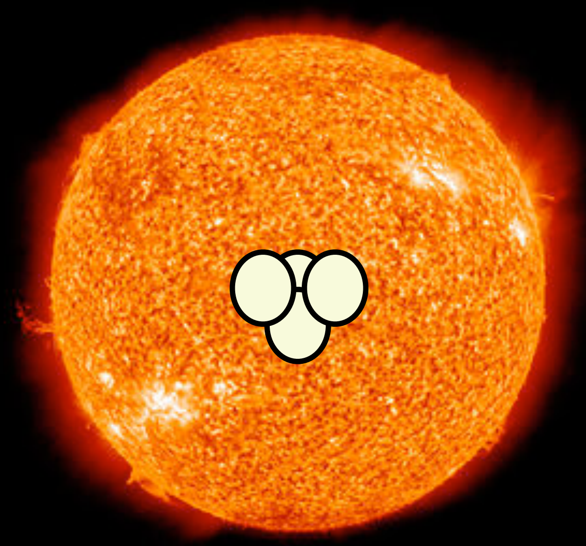


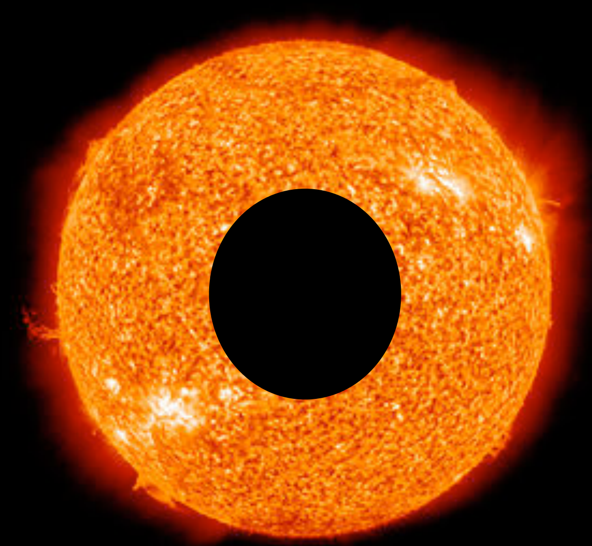


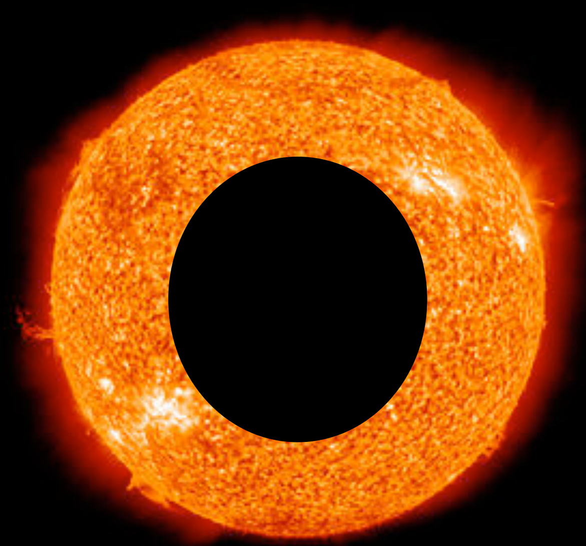


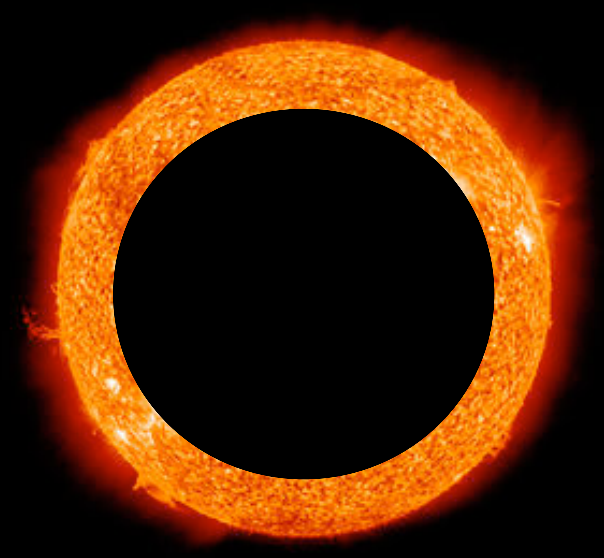












Bosonic Asymmetric Dark Matter

No Fermi pressure but Heisenberg uncertainty keeps bosons from collapse

$$\frac{GNm^2}{r} \simeq \frac{\hbar}{r} \longrightarrow M_{crit} = \frac{2M_{\rm Pl}^2}{\pi m}\sqrt{1+\frac{M_{\rm Pl}^2}{4\sqrt{\pi}m}\sigma^{1/2}}$$

$$\sigma = \lambda^2/(64\pi m^2)$$

repulsive interactions

BEC accelerates collapse

$$T_c = \left(\frac{n}{\zeta(3/2)}\right)^{2/3}\frac{2\pi\hbar^2}{mk_B} \approx 3.31\,\frac{\hbar^2 n^{2/3}}{mk_B}$$

$$N_{\rm BEC} \simeq 2\times 10^{36}$$

$$r_{\rm th} \simeq 2\ {\rm m}\left(\frac{T_c}{10^5{\rm K}}\right)^{1/2}\left(\frac{m}{\rm GeV}\right)^{-1/2} \longrightarrow r_c = \left(\frac{8\pi}{3}G\rho_c m^2\right)^{-1/4} \simeq 1.6\times 10^{-4}\left(\frac{\rm GeV}{m}\right)^{1/2}{\rm cm}$$

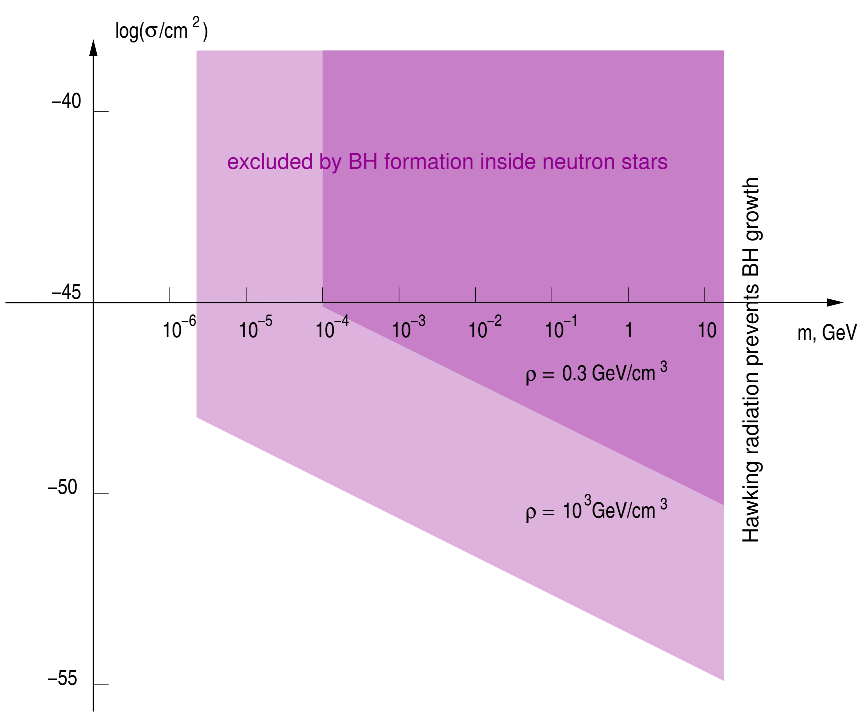
Evolution of the Black Hole

$$\frac{dM}{dt} = \frac{4\pi\rho_c G^2 M^2}{c_s^3} - \frac{1}{15360\pi G^2 M^2}$$

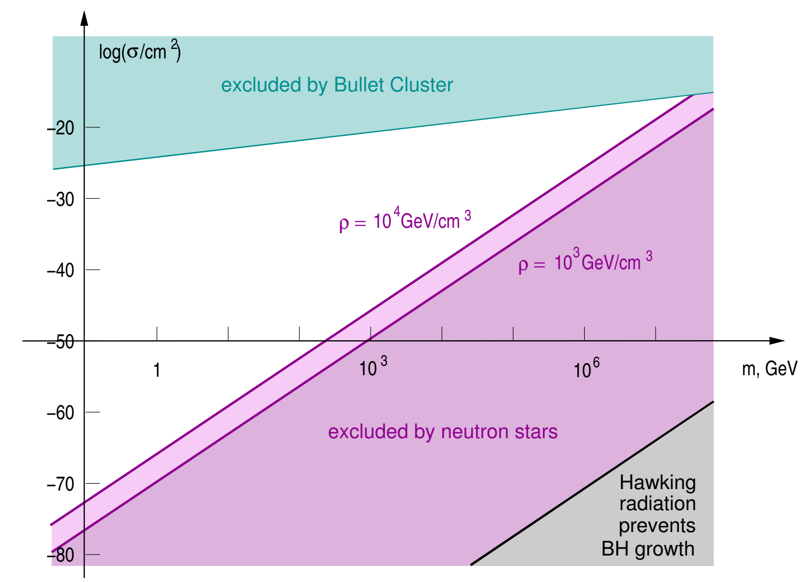
Bondi accretion

Hawking Radiation

Bosonic Asymmetric Dark Matter



CK, Tinyakov PRL '11



Self-Interacting DM

If WIMP is a composite of fermions

$$\Lambda_{crit} = m^{1/3} M_{\rm Pl}^{2/3} \left(1 + \frac{\lambda m_{pl}^2}{32\pi m^2}\right)^{-1/3}$$

If WIMP is a composite of fermions above that scale, the bosonic constraints still hold

The effect of Rotation I

For the scenario to be realised one should make sure that all conditions are met!

The accretion is never perfectly spherical because the neutron star rotates usually with high frequencies.

Can rotation slow down the accretion to the point that invalidate the constraints?

The conditions for Bondi accretion are valid as long as the angular momentum of an infalling piece of matter is much smaller than the keplerian one in the last stable orbit

The mass of the black hole must be larger than

$$M_{\rm crit} = \frac{1}{12^{3/2}} \left(\frac{3}{4\pi \rho_c} \right)^2 \left(\frac{\omega_0}{G} \right)^3 \frac{1}{\psi^3}$$

$$M_{\rm crit} = 2.2 \times 10^{46} P_1^{-3} \text{ GeV}$$

CK, Tinyakov '13

viscosity of nuclear matter can help!

It subtracts angular momentum at the initial stage where the black hole is still small

in the final stages Bondi accretion is not valid but the star is seconds away from destruction!

The effect of Rotation II

A maximally spinning black hole will stop the accretion

$$a = J/GM^2$$

$$\frac{1}{a}\frac{da}{dt} = \frac{1}{J}\frac{dJ}{dt} - 2\frac{1}{M}\frac{dM}{dt}$$

$$\frac{1}{a}\frac{da}{dt} = \frac{1}{J}\omega_0 r_s^2 \frac{dM}{dt} - \frac{g(a)}{G^2M^3} - \frac{2}{M}\frac{dM}{dt}$$

After formation the black hole spins down, then it spins up and at the last stages it spins down again

$$a_{\rm max} = 2 \times 10^{-23} T_5^4 / P_1^{10}$$

Temperature Considerations

Radiation from in falling matter can in principle impede further accretion in two ways:

Reduce viscosity

Increase radiation pressure

e-e Bremsstrahlung close to the horizon is the dominant radiation mechanism

$$\epsilon = \frac{L_{ee}}{dM/dt} \simeq 5 \times 10^{-12} T_5 \left(\frac{M}{M_0}\right)$$

$$\delta T = \frac{L_{ee}}{4\pi k r} \simeq 458 \left(\frac{M}{M_0}\right)^2 \left(\frac{r_B}{r}\right) \mathrm{K}$$

Self-Interacting Dark Matter

"Chandrashekhar Limit for WIMPs"

$$\frac{GNm^2}{r} > k_F = \left(\frac{3\pi^2 N}{V}\right) = \left(\frac{9\pi}{4}\right)^{1/3} \frac{N^{1/3}}{r}$$

N=10^57/m^3!!!

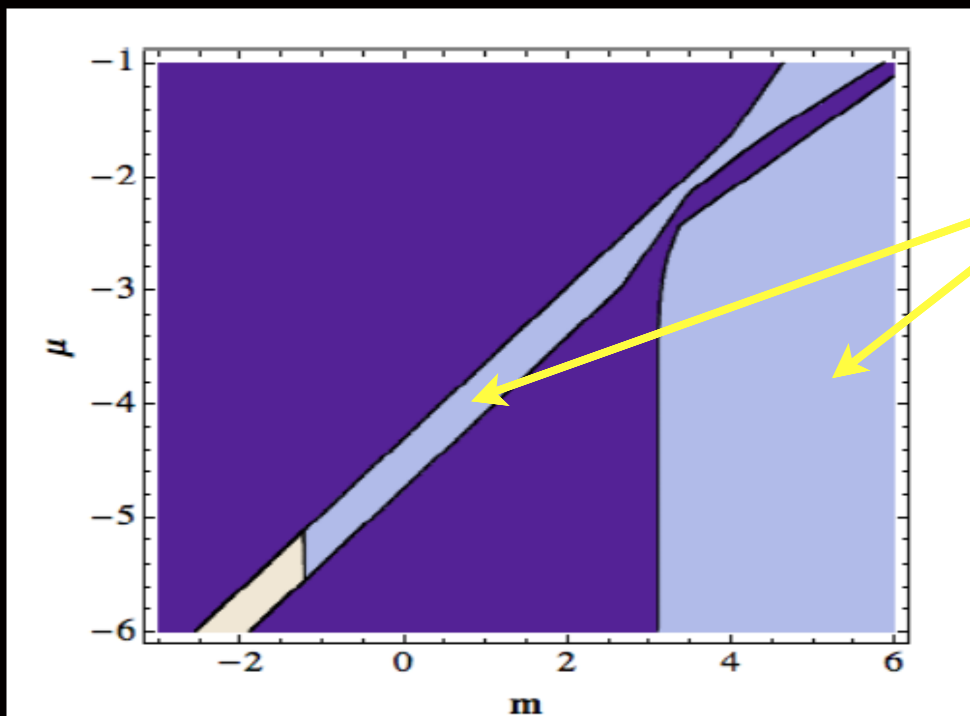
Yukawa-type WIMP self-interactions can explain the flatness of dwarf galaxies Spergel-Steinhardt '99, Loeb-Weiner '11

$$\alpha\phi\bar{\psi}\psi$$

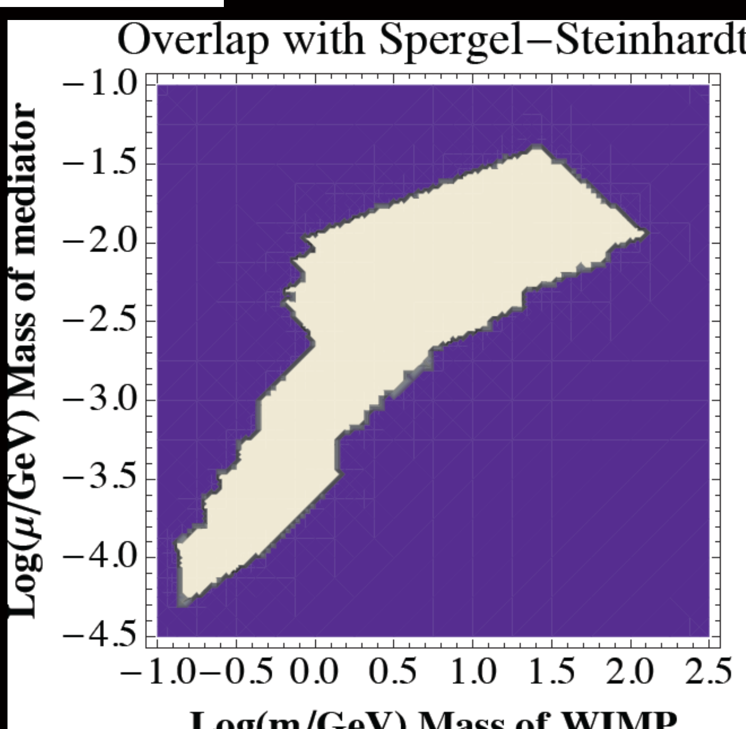
$$V(r) = -\alpha \exp[-\mu r]/r$$

Yukawa self-interactions can alleviate the effect of the Fermi pressure, leading to a gravitational collapse with dramatically lower amount of captured WIMPs

Self-Interacting Dark Matter

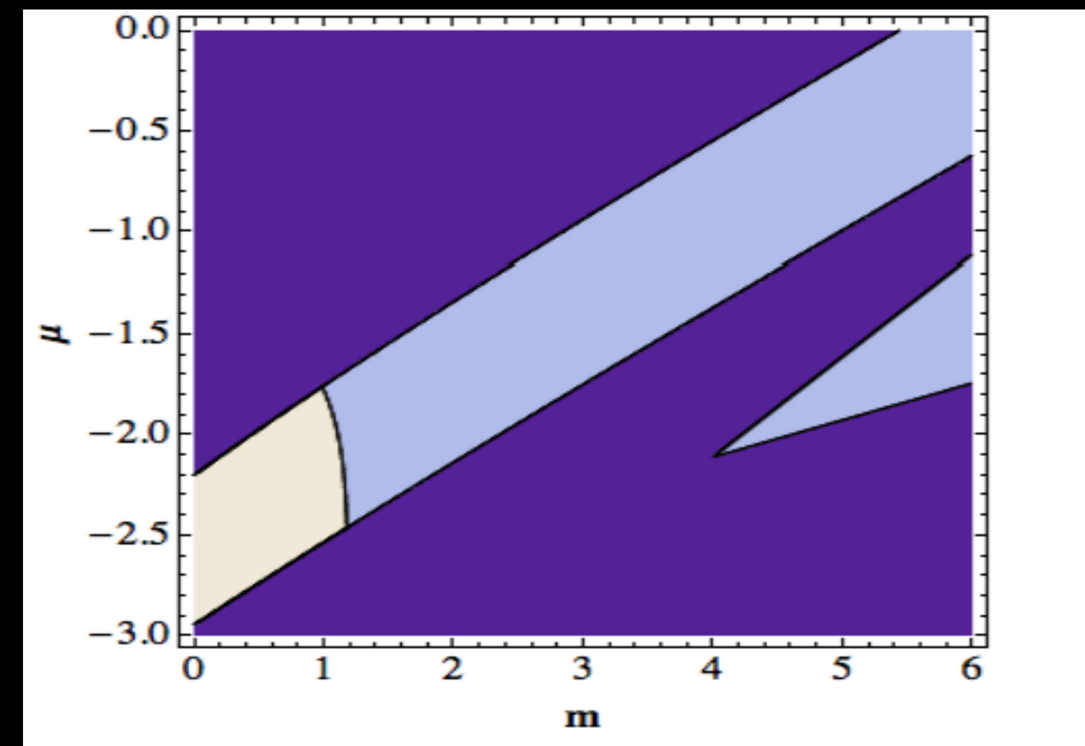


$\alpha = 10^{-5}$



Exclusion regions

CK PRL'12



$\alpha = 0.1$

Loeb-Weiner

$$(m_\chi/10\mathrm{GeV})(m_\phi/100\mathrm{MeV})^2 \sim 1$$

Pulsar Spindown

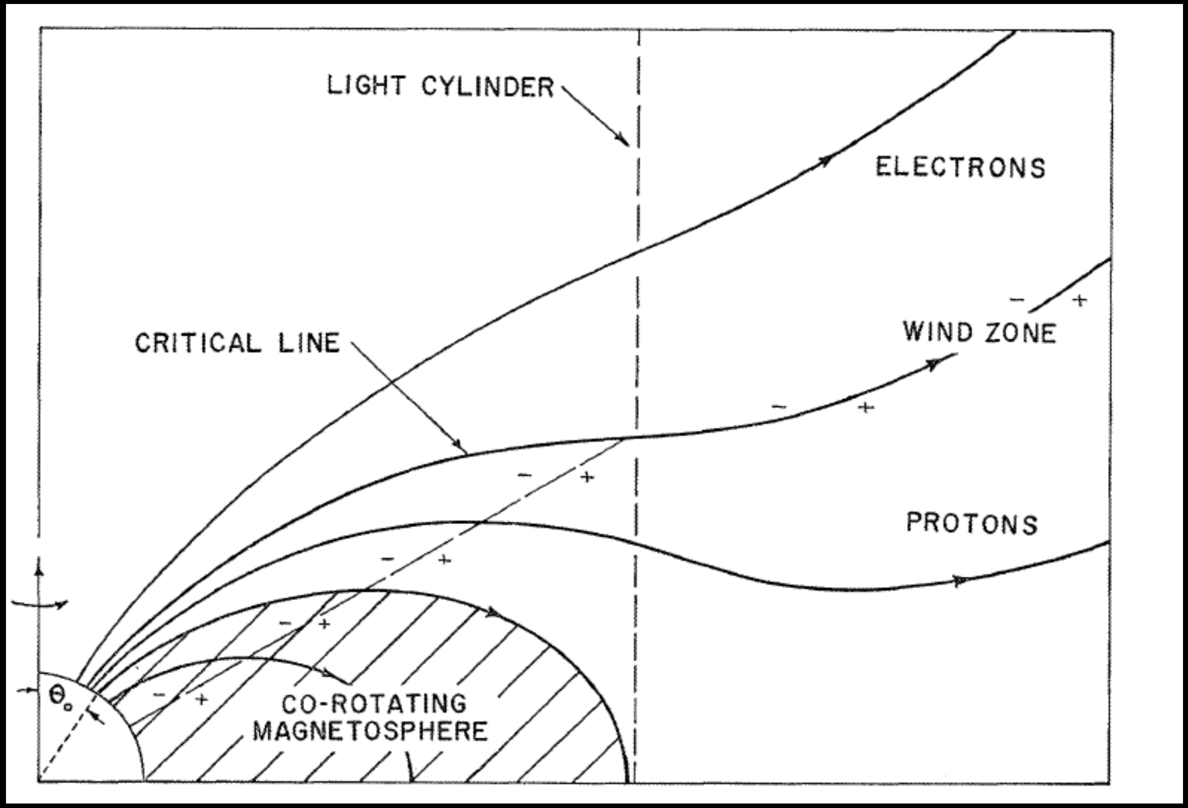
Pulsars spin down according to 2 basic mechanisms

- Magnetic Dipole Radiation
- Aligned Rotator

$$L = L_{\rm orth} \sin^2\theta + L_{\rm align}\cos^2\theta$$

$$L_{\rm orth} = \frac{B_0^2\Omega^4 R^6}{4c^3}, \qquad L_{\rm align} = \frac{B_0 \Omega \Omega_F R^3 I}{2c^2}$$

Contopoulos, Spitkovsky '06



$$I = I_{\rm GJ} = \frac{B_0 R^3 \Omega^2}{2c} \simeq 1.4\times 10^{30} e \left(\frac{B_0}{10^{12}\,G}\right)\left(\frac{s}{P}\right)^2 \ {\rm s}^{-1}$$

$$\dot{E} = \mathcal{I}\Omega\dot{\Omega} = \tfrac{2}{5}MR^2\Omega\dot{\Omega} = -L$$

Julian Goldreich '69

Pulsar Spindown

Can accretion of millicharged particles affect the spinning?

At steady state a NS should expel as much charge as it accumulates

$$\dot{\Omega} = -\frac{5B_0^2\Omega^3R^4}{8Mc^3}\left[\sin^2\theta + \left(1-\frac{\Omega_{\rm death}}{\Omega}\right)\left(1+\frac{I_{\rm DM}}{I_{\rm GJ}}\right)\cos^2\theta\right]$$

Observable independent of the magnetic field braking index

$$n = \ddot{\Omega}\Omega/\dot{\Omega}^2$$

$$n = 3 - \frac{2\lambda\cos^2\theta}{1+\lambda\cos^2\theta}$$

$$\lambda = \frac{I_{\rm DM}}{I_{\rm GJ}}$$

If the DM current is similar to the GJ current, the braking index can significantly different than 3

Pulsar Spindown

TABLE 2

SPIN AND INFERRED PARAMETERS FOR PULSARS WITH MEASURED n, ORDERED BY SPIN-DOWN AGE

Pulsar	n^{a}	ν (s^{-1})	$\dot{\nu}$ (10^{-11} s^{-2})	τ_c [b] (yr)	τ [c] (yr)	B_{di} [d] (10^{12} G)	$\dot{E}$ [e] (10^{36} ergs s^{-1})	Reference
J1846−0258	2.65(1)	3.07	−6.68	723	884	49	8.1	1
B0531+21	2.51(1)	30.2	−38.6	1240	1640	3.8	460	2
B1509−58	2.839(3)	6.63	−6.76	1550	1690	15	18	3
J1119−6127	2.91(5)	2.45	−2.42	1610	1680	42	2.3	4
B0540−69	2.140(9)	19.8	−18.8	1670	2940	5.1	150	5
B0833−45[f]	1.4(2)	11.2	−1.57	11300	57000	3.4	6.9	6

[a] Uncertainties on n are in the last digit.
[b] Characteristic age is given by $\tau_c = \nu/2\dot{\nu}$.
[c] Inferred upper-limit timing age given n, $\tau = \nu/(n-1)\dot{\nu}$, assuming $\nu_i \gg \nu$.
[d] Dipole magnetic field estimated by $B_{di} = 3.2\times10^{19}(-\dot{\nu}/\nu^3)$ G, assuming $\alpha = 90°$ and $n = 3$.
[e] Spin-down luminosity, $\dot{E} \equiv 4\pi^2 I\nu\dot{\nu}$, where it is assumed that $I = 10^{45}$ g cm^2 for all pulsars.
[f] Braking index for the Vela pulsar was not determined from a standard timing analysis due to the large glitches experienced by this object. Instead, measurements of $\dot{\nu}$ were obtained from assumed "points of stability" 100 days after each glitch (see Lyne et al. 1996 for details).

REFERENCES.—(1) This work; (2) Lyne et al. 1993; (3) Livingstone et al. 2005b; (4) Camilo et al. 2000; (5) Livingstone et al. 2005a; (6) Lyne et al. 1996.

Dark Matter accretion

Electromagnetic

$$F_{\rm I} \simeq 1.0 \times 10^{29} \left(\frac{\rho_{\rm DM}}{0.3\,{\rm GeVcm^{-3}}}\right)\left(\frac{\rm 1GeV}{m}\right)\left(\frac{P}{\rm s}\right)\ {\rm s}^{-1}$$

Gravitational

$$F_{\rm II} = 4.2 \times 10^{26} \left(\frac{\rho_{\rm DM}}{0.3\,{\rm GeVcm^{-3}}}\right)\left(\frac{\rm 1GeV}{m}\right) f\ {\rm s}^{-1}$$

Condition

$$r_L << R_{\rm curv}$$

$$r_L = \frac{m v_\perp}{\epsilon e B}$$

$$R_{\rm curv}(r,\gamma) = \left|\frac{1}{B}\frac{dB}{dr}\right|^{-1}$$

Pulsar Spindown

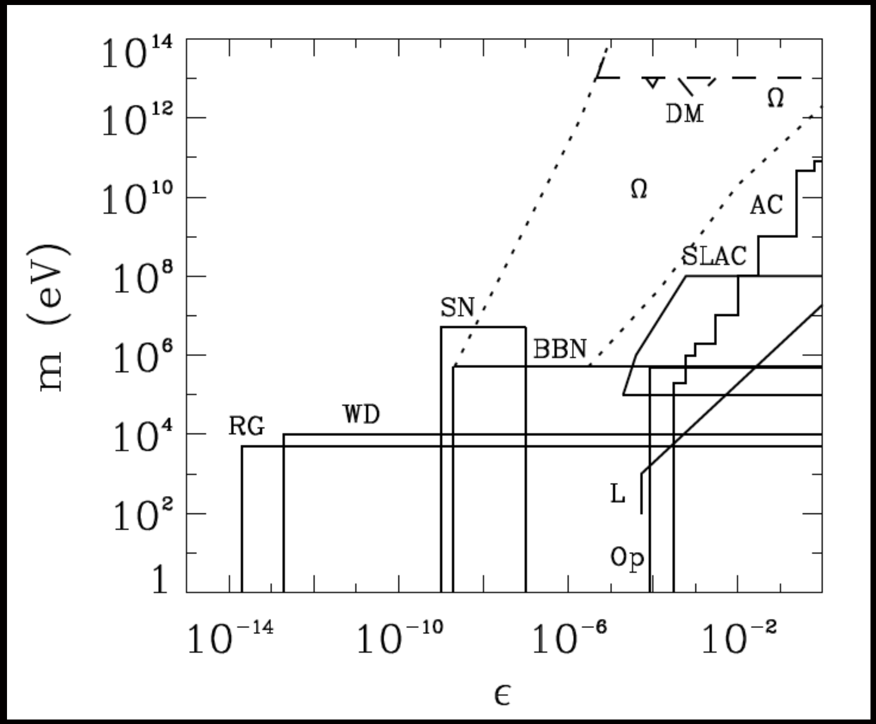
Electromagnetic

$$\epsilon \left(\frac{\mathrm{GeV}}{m}\right) \simeq 14 \left(\frac{0.3\,\mathrm{GeV cm^{-3}}}{\rho_{\mathrm{DM}}}\right) \left(\frac{\mathrm{s}}{P}\right)^3 \left(\frac{B_0}{10^{12}\,G}\right)$$

CK, Perez-Garcia '14

Gravitational

$$\epsilon \left(\frac{\mathrm{GeV}}{m}\right) \simeq \frac{3.3 \times 10^3}{f} \left(\frac{0.3\,\mathrm{GeV cm^{-3}}}{\rho_{\mathrm{DM}}}\right) \left(\frac{\mathrm{s}}{P}\right)^2 \left(\frac{B_0}{10^{12}\,G}\right)$$



Davidson,Hannestad, Raffelt '00

$$\epsilon^2 \gtrsim 5 \cdot 10^{-11}\,\mathrm{GeV}^{-1/2} \frac{m_X}{\sqrt{\mu_{X,e}} + \sqrt{\mu_{X,p}}}$$

Dolgov,Dubovsky,Rubtsov,Tkatchev '13

Can be satisfied for example for MeV, ε~10^-6 and either higher DM density or lower magnetic field

Stopping WIMPs

Different mechanisms of slowing down WIMPs:

- Nuclear Stopping
- Conductor Stopping
- Insulator Stopping

WIMPs don't have enough energy to ionise atoms, Bohr-Bethe formula is not applicable

Conductor Stopping
Free Fermi gas approximation

$$\frac{dE}{dx} = -\frac{4\varepsilon^2 e^4 m_e^2}{3\pi} v C_1(\chi)$$

$$C_1(\chi) = \int_0^1 \frac{z^3 dz}{(z^2 + \chi^2 f_1(0,z))^2}$$

$$v_F = (3\pi^2 n)^{1/3}/m_e$$

$$\chi^2 = e^2/(\pi v_F)$$

$$f_1(0,z) = \frac{1}{2} + \frac{1}{4z}(1 - z^2)\ln\left|\frac{z+1}{z-1}\right|$$

$$\sqrt{E_{\mathrm{in}}} - \sqrt{E_{\mathrm{f}}} = \frac{2\sqrt{2}\varepsilon^2 e^4 m_e^2}{3\pi\sqrt{m_X}} C_1(\chi) L$$

Insulator Stopping

$$\frac{dE}{dx} = -\frac{4\pi\varepsilon^2 e^4}{m_e v^2} nZ\mathcal{L}$$

$$2m_e v^2/\hbar\omega$$

$$E_{in}^2 - E_f^2 = (4\pi)^3 \frac{m_X}{m_e}\varepsilon^2\alpha^2 nZ\overline{\mathcal{L}}L$$

Stopping WIMPs

Different mechanisms of slowing down WIMPs:

- Nuclear Stopping
- Conductor Stopping
- Insulator Stopping

Nuclear Stopping

$$\frac{dE}{dx} = -n_N \int_{E_R^{min}}^{E_R^{max}} \frac{d\sigma}{dE_R} E_R dE_R$$

$$E_R^{max} = 4 m_X m_N E / (m_X + m_N)^2$$

$$E_R^{min} = 1/(2 m_N a_0^2)$$

Millicharged & Light mediators

$$\frac{d\sigma}{dE_R} = \frac{2\pi\alpha_{EM}^2 \varepsilon^2 Z^2}{m_N v^2 E_R^2}$$

$$\mathrm{Ei}\left(2 \ln bE_{in}\right) - \mathrm{Ei}\left(2 \ln bE_f\right) = \frac{\pi n_N \alpha_{EM}^2 \varepsilon^2 Z^2 b^2 m_X L}{m_N}$$

Contact Interactions

$$\frac{d\sigma}{dE_R} = \frac{m_N \sigma_N}{2\mu_N^2 v^2} = \frac{m_N \sigma_n A^2}{2\mu_p^2 v^2}$$

$$\ln \frac{E_{in}}{E_f} = \frac{2 n_N \sigma_p A^2 \mu_N^4 L}{m_X m_N \mu_p^2}$$

$$\mathrm{Ei}(x) \equiv - \int_{-x}^{\infty} \frac{e^{-t}}{t} dt$$

$$b \equiv \frac{8 m_N^2 m_X a_0^2}{(m_N + m_X)^2}$$

Diurnal Modulations & Light DM

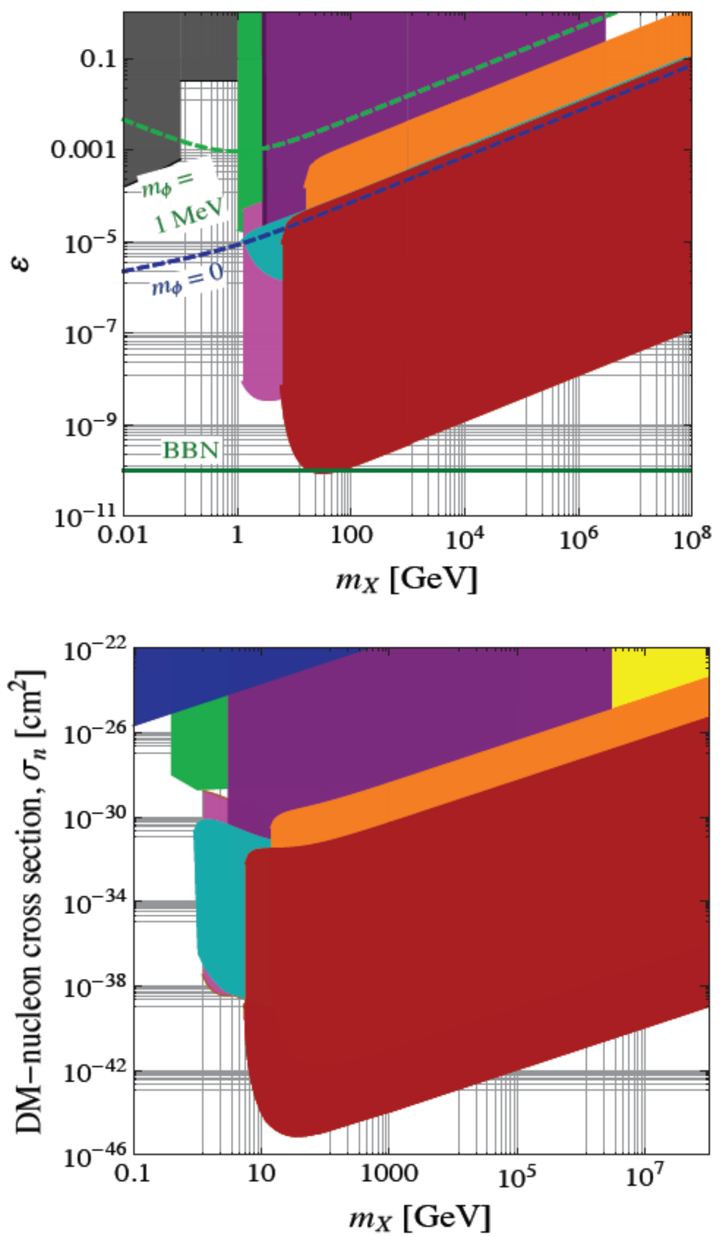


FIG. 2: The allowed ($\varepsilon - m_X$) parameter space for millicharged and light mediators in the top panel, and contact interactions in the ($\sigma_n - m_X$) plane in the bottom panel based on LUX (red), CRESST-1 (cyan), RRS balloon (purple), DAMIC [18] (magenta), XQC [23] (green), CDMS-1 (orange), and the SLAC accelerator limits [21, 22] on millicharged particles (gray). In the contact interaction case, we also include CMB limits [26] (blue), and those from the IMP8 experiment [6, 29] (yellow) constraints.

There is phase space that will not be covered by underground experiments even if they lower their energy threshold due to effective stopping by the rock.

Need for detectors in shallow sites or on surface

However this would increase the background!

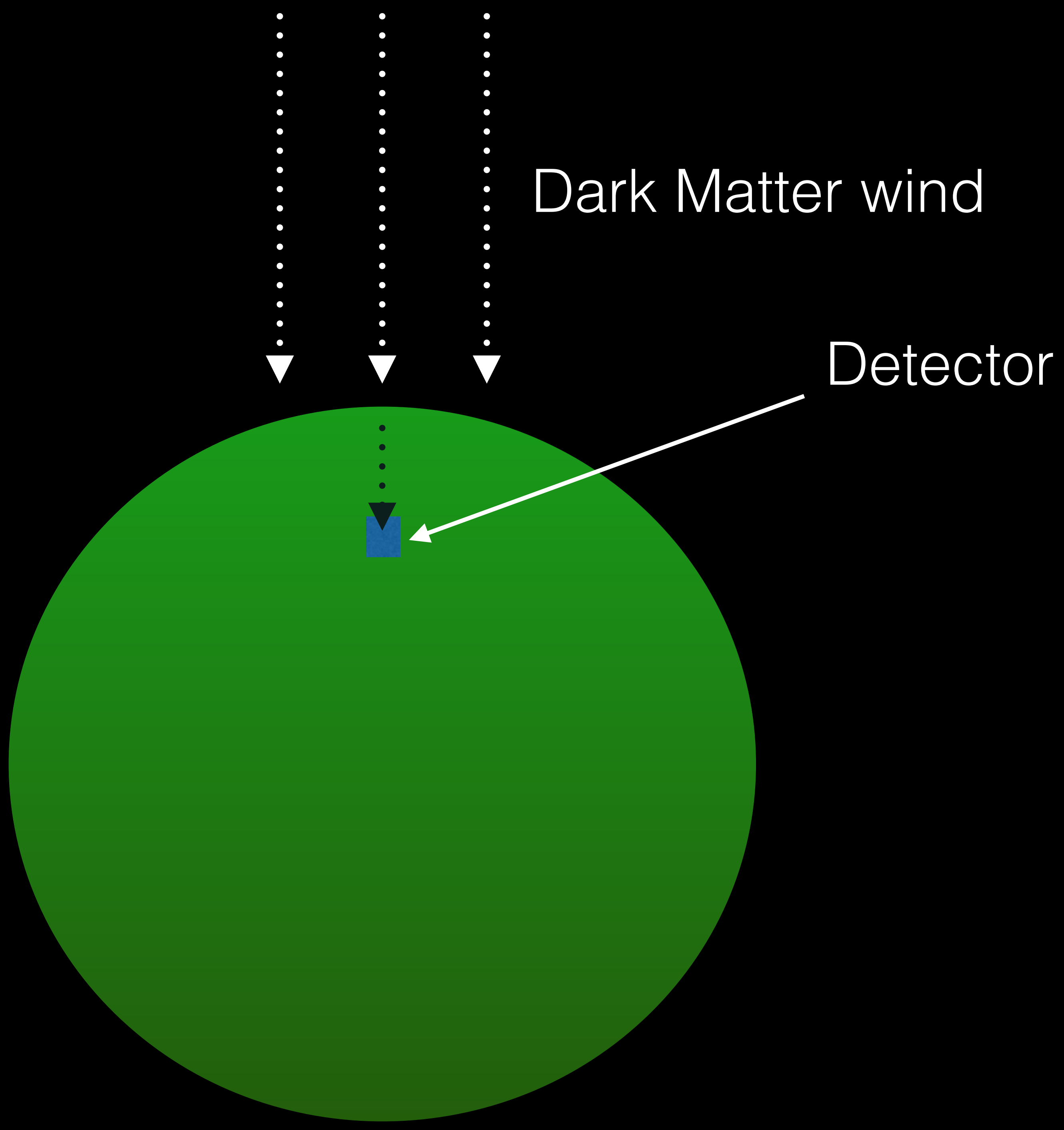
Diurnal Modulation

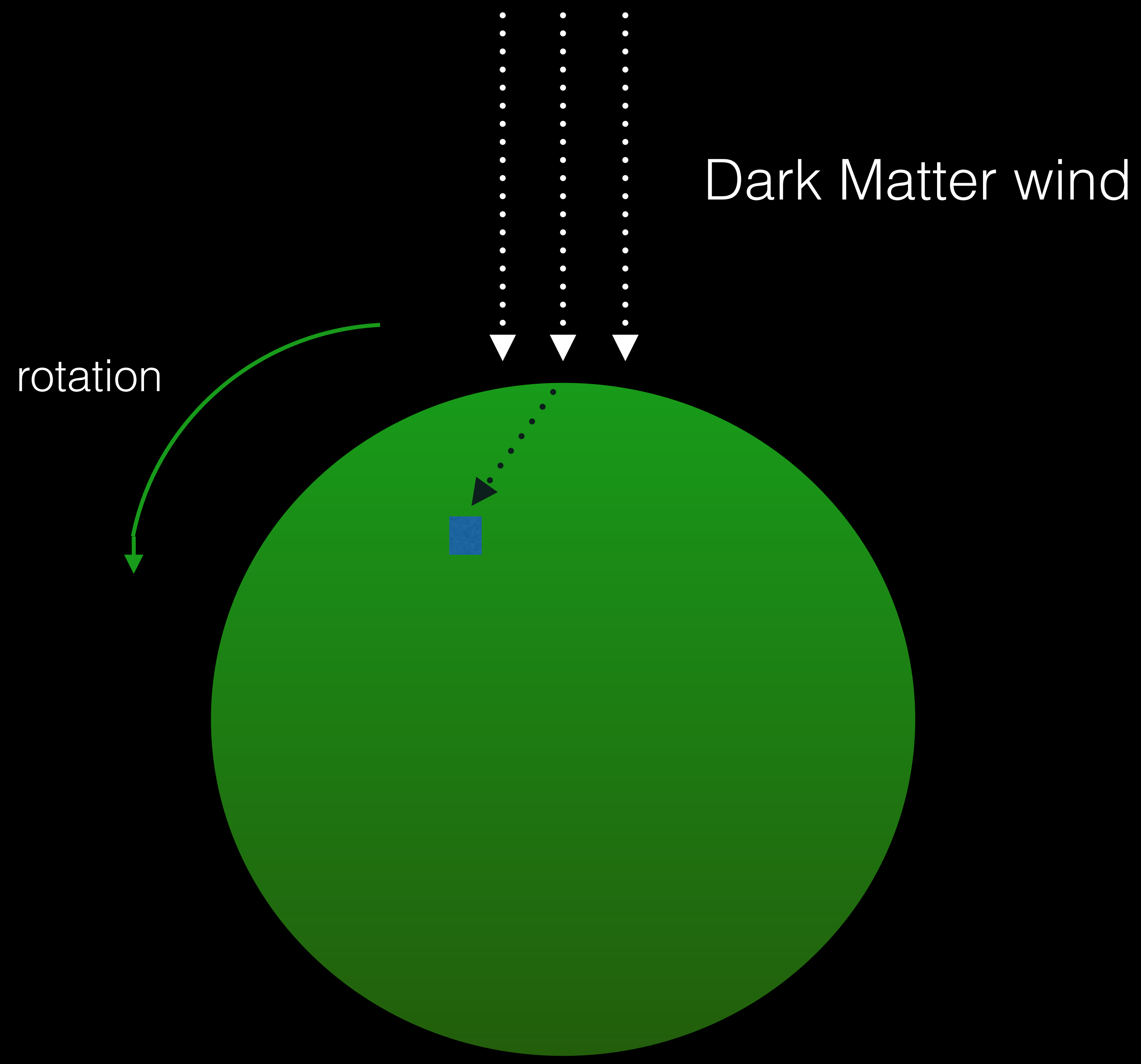
DM particles can reach from top but not from the bottom

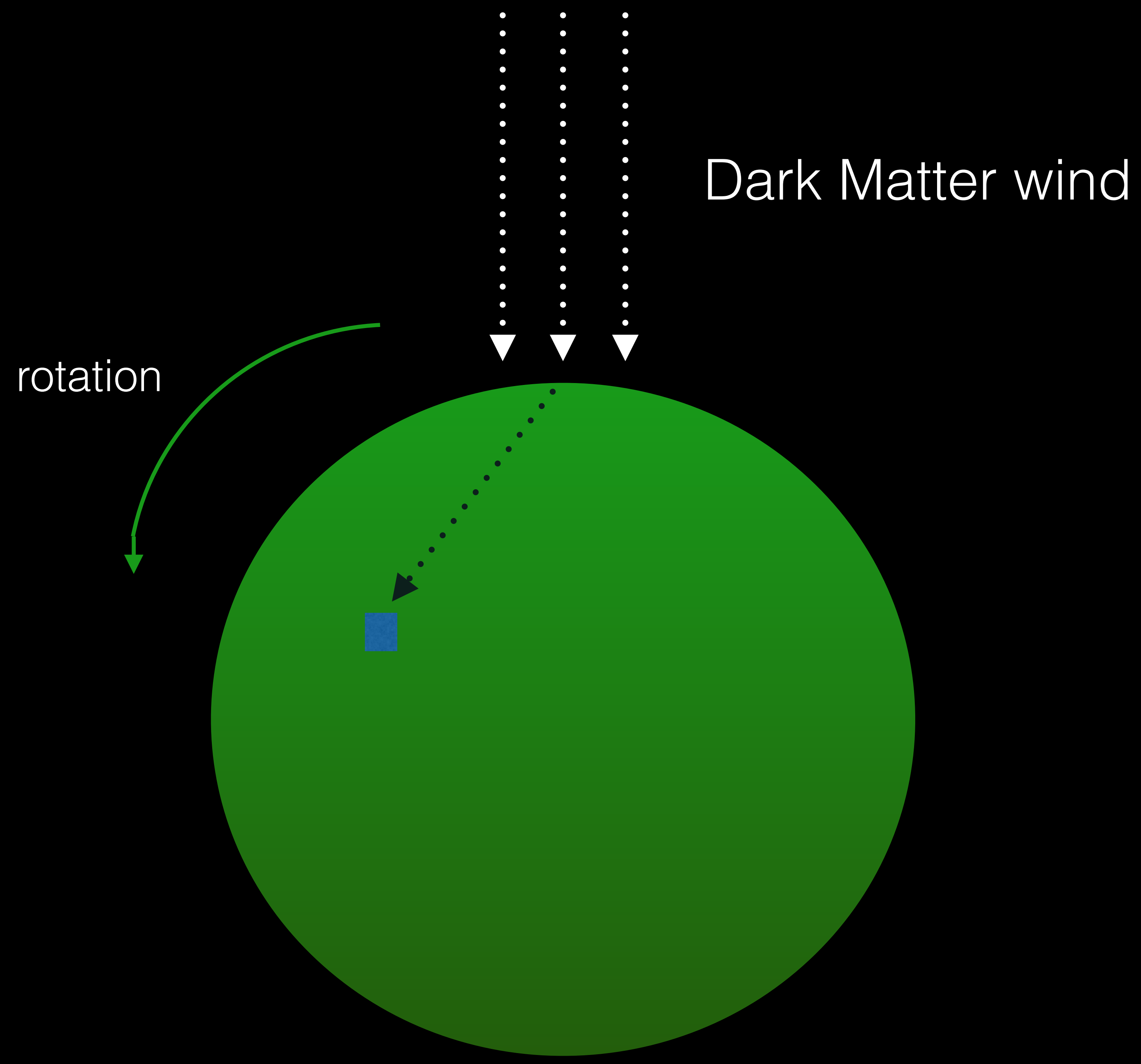
This gives a signal that modulates on a daily basis

Ideal for portable experiments like DAMIC

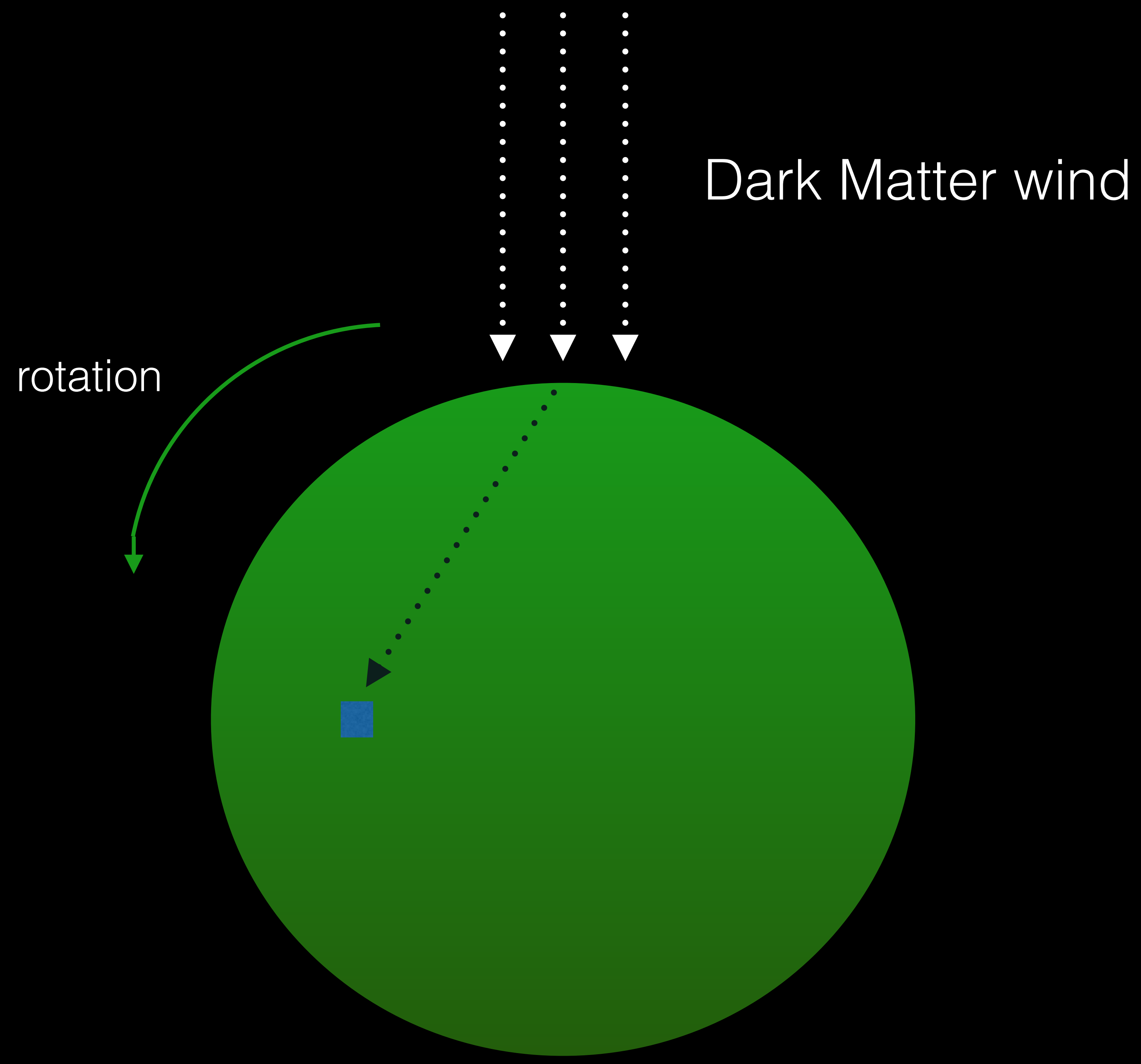
CK, Shoemaker '14

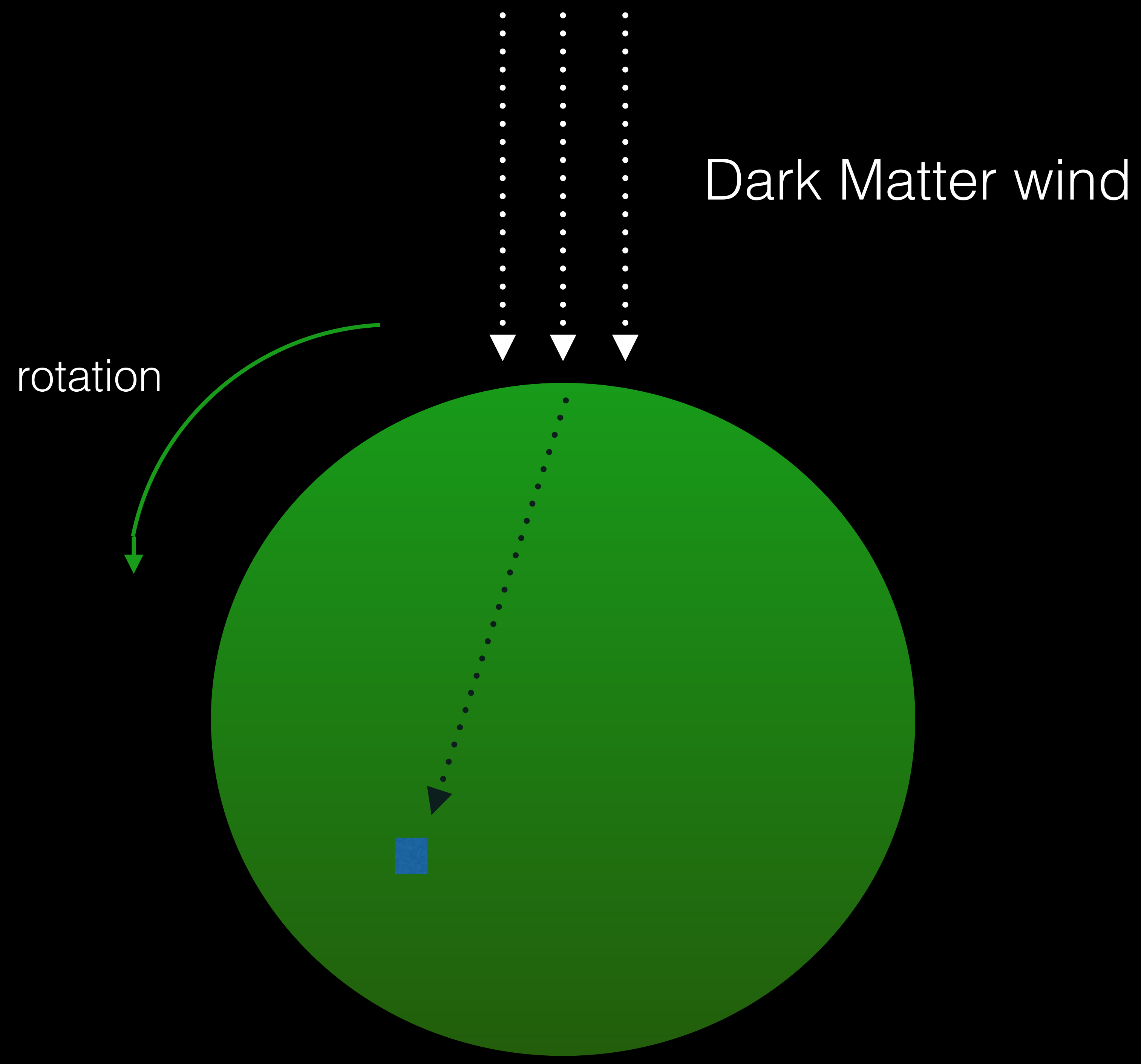

Dark Matter wind
Detector

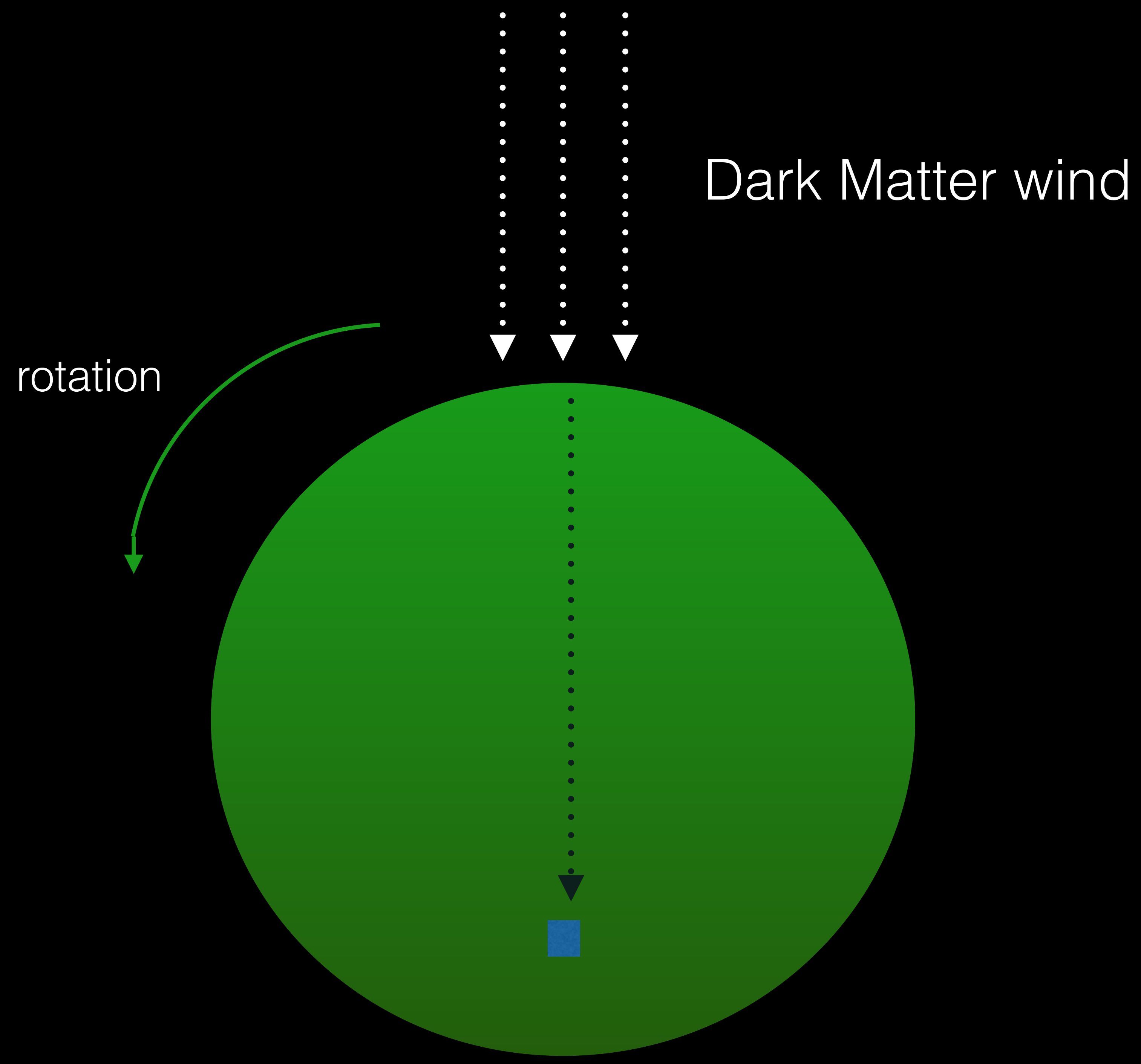

Dark Matter wind
rotation



Dark Matter wind
rotation


Dark Matter wind
rotation


Dark Matter wind
rotation


Dark Matter wind
rotation

Estimating the Asymmetry

$$f(v)d^3v = \left(\frac{1}{\pi v_0^2}\right)^{3/2} e^{-\frac{v^2+v_E^2}{v_0^2}} e^{-\frac{2vv_E}{v_0^2}\cos\delta} d^3v$$

δ is the angle between $\vec{v}$ and $\vec{v}_E$

length traveled underground

$$\ell(\psi) = (R_\oplus - \ell_D)\cos\psi + \sqrt{(R_\oplus - \ell_D)^2\cos^2\psi - (\ell_D^2 - 2R_\oplus \ell_D)}$$

ψ to be the angle between $\vec{v}$ and $\hat{n}$

$$\cos\psi = \hat{v}\cdot\hat{n} = \cos\theta_l \cos\omega t \cos\theta \cos\phi + \cos\theta_l \sin\omega t \cos\theta \sin\phi \pm \sin\theta_l \sin\theta.$$

$$\frac{dR}{dE_R} = \frac{\rho_X}{m_X m_N}\int_{v_{min}}^{v_{esc}} \frac{d\sigma}{dE}(E', E_R) v^3 f(v) dv d\cos\theta d\phi$$

How the Asymmetry can be detected?

- Portable detectors like DAMIC can be placed either on the surface or in shallow sites
- The best latitude for most candidates is $\theta_l \simeq -\alpha$ where α is the angle between the earth's rotation axis and earth's velocity ~ 43 degrees
- Best locations at the south hemisphere e.g. Chile, Argentina, New Zealand
- The signal varies during the year due to the change in the velocity of the earth with respect to the rest frame of the dark matter halo.

The Dark Side of the Stars

Compact stars can reveal a lot of information about the nature of DM putting constraints on its properties complementary to direct searches.

- Observation of cold neutron stars can exclude thermally produced dark matter.
- Asymmetric dark matter:

1. keV to few GeV non-interacting bosonic dark matter is excluded.
2. Part of fermionic WIMP self-interactions excluded.
3. Constraints on WIMP-nucleon spin-dependent interactions.

- Millicharged dark matter could slow down pulsars faster than other mechanisms predict.

Diurnal Modulated Signal from WIMP-nucleus stopping as DM travels underground