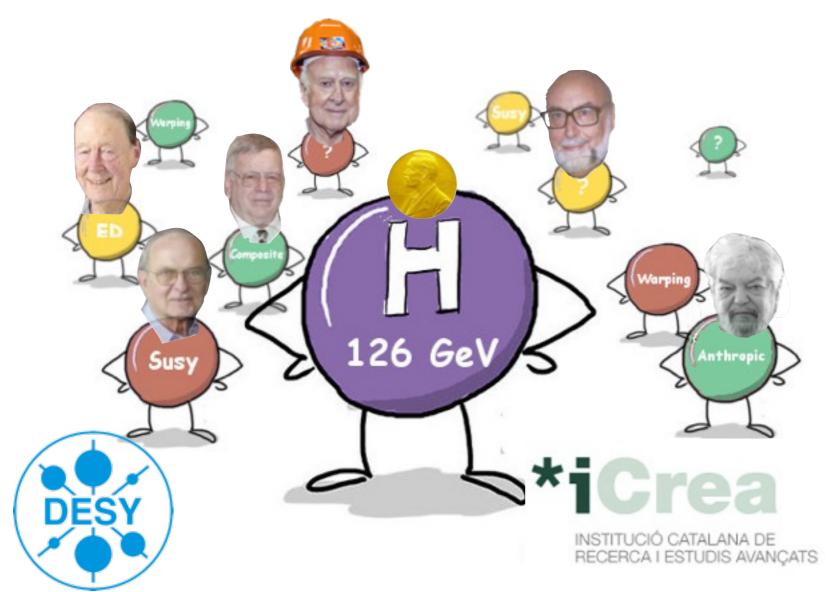
Beyond the Standard Model

CERN summer student lectures 2015



Lecture 1/5

Christophe Grojean

DESY (Hamburg) ICREA@IFAE (Barcelona) (christophe.grojean@cern.ch)

What is physics beyond the Standard Model?



I don't know. Nobody knows If it were known, it would be part of the SM! You won't learn during these lectures what is BSM you'll learn what BSM could be "Looking and not finding is different than not looking"

we'll study the limitations/defaults of the SM as a guide towards BSM

Outline

Monday

O general introduction, units

• Tuesday

• Higgs physics as a door to BSM

Wednesday

• Naturalness: small and large numbers in a quantum world

Thursday

O grand unification, proton decay

• supersymmetry

0 extra dimensions

□Friday

• cosmological interplay

Recommended Readings

popular account

0 "The Zeptospace odyssey" by Gian-Francesco Giudice <u>CERN library link</u>

\Box fun physics

O "Order-of-magnitude physics" by S. Mahajan, S. Phinney and P. Goldreich <u>available for free online</u>

technical accounts

O "Journeys beyond the Standard Model" by P. Ramond <u>CERN library link</u>

(The elementary particles

Particle physics is special

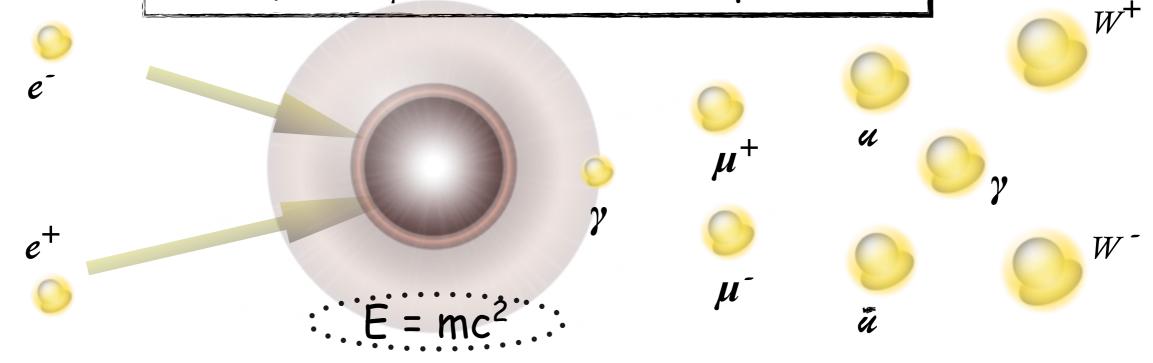
• Chemistry: reorganization of matter

the various constituents of matter reorganize themselves in different structures

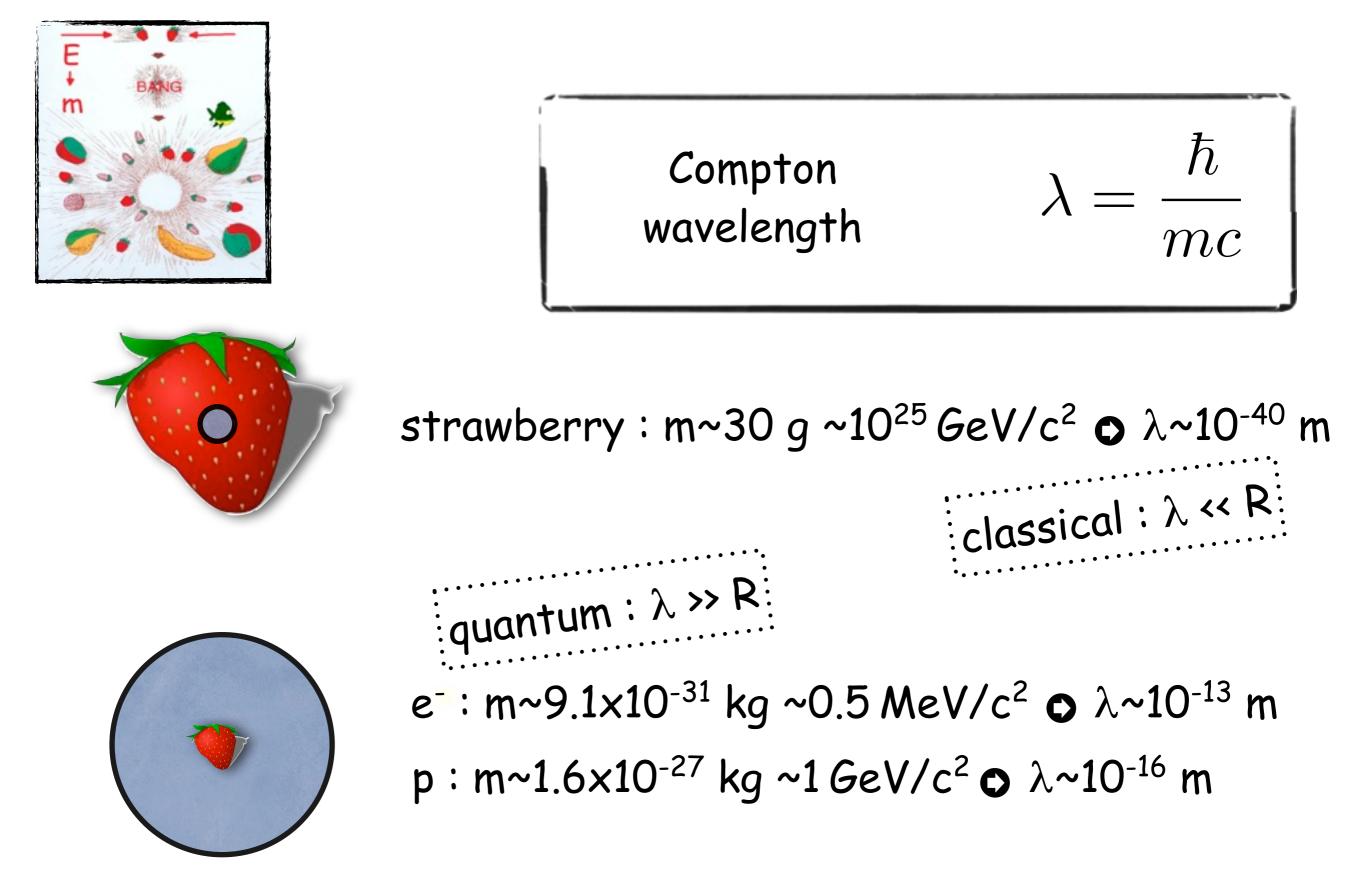
$CH_4 + 2 O_2 \rightarrow CO_2 + 2 H_2O$

• Particle physics: transformation matter \Leftrightarrow energy

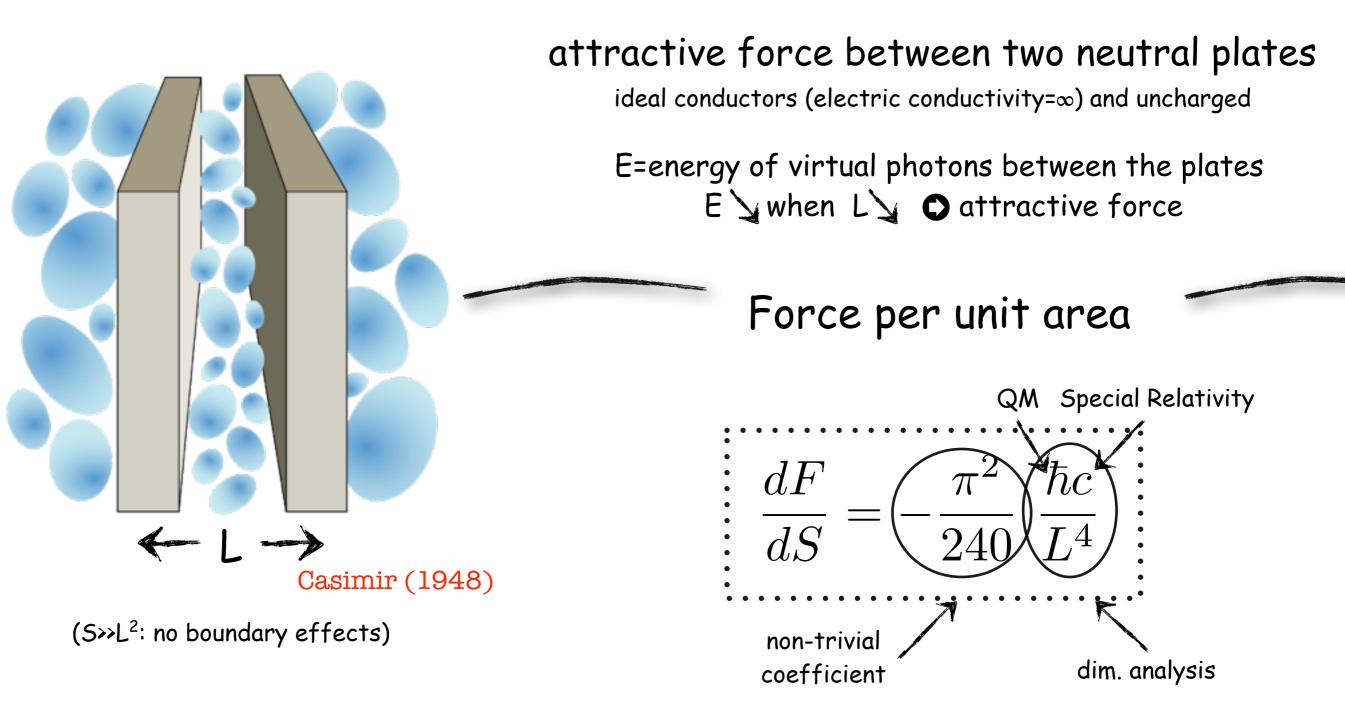
NB: Q, L_e, L_{μ}, B = conserved quantities



Classical vs. Quantum Collisions



Vacuum fluctuations



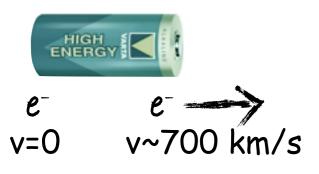
numerically: pressure of ~ 1 atm for a 10nm separation

8

The quantum vacuum is not empty

BSM

Energy Scales of Particle Physics



1 TeV = 10¹² eV 1eV=energy of an electron accelerated by a potential difference of 1 volt 1 eV= 1.6×10⁻¹⁹ Joule



1 kg sugar = 4000 kCalories= 17 millions of Joule but 1 kg sugar ≈ 10²⁷ protons 0.1 eV / protons

If one wanted to accelerate each protons contained in 1kg of sugar to 14 TeV, (s)he would need the caloric energy contained in 10¹⁴kg of sugar* or 1% of the total energy produced yearly

*yearly worldwide production of sugar=150 millions of tons \approx 10¹¹ kg

Classical/Quantum EM & Antimatter

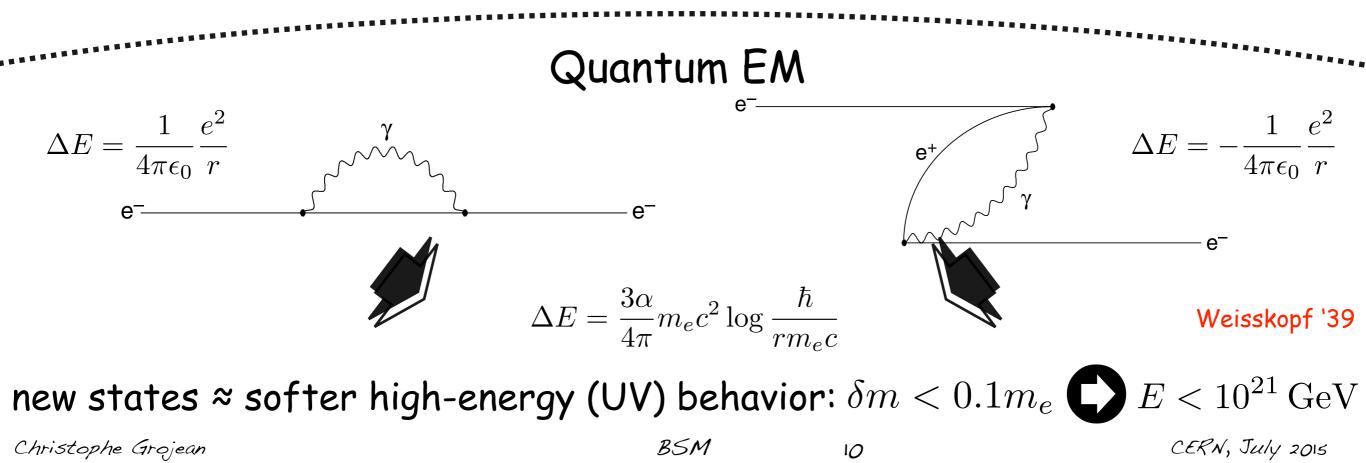
an electron makes an electric field which carries an energy

$$\Delta E_{\rm Coulomb}(r) = \frac{1}{4\pi\epsilon_0} \frac{e^2}{r}$$

and interacts back to the electron and contributes to its mass $\delta mc^2 = \Delta E$

$$\delta m < m_e$$
 (C) $r > r_e \equiv \frac{e^2}{4\pi\epsilon_0 m_e c^2} \sim 10^{-13} \text{ m}$ i.e. $E < \frac{\hbar c}{r_e} \sim 5 \text{ MeV}$

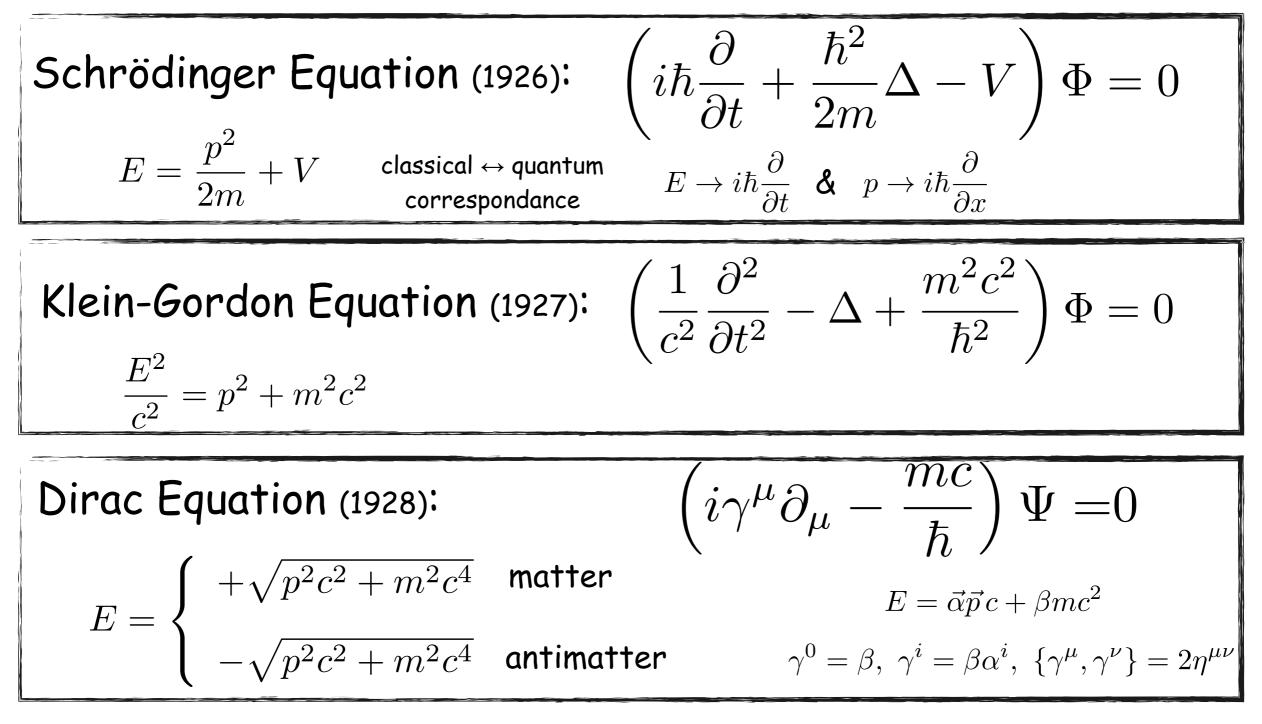
At shortest distances or larger energies, classical EM breaks down



Antimatter and Dirac equation

Schrödinger's equation (1926) is non-relativistic

(cannot account for creation/annihilation of particles)

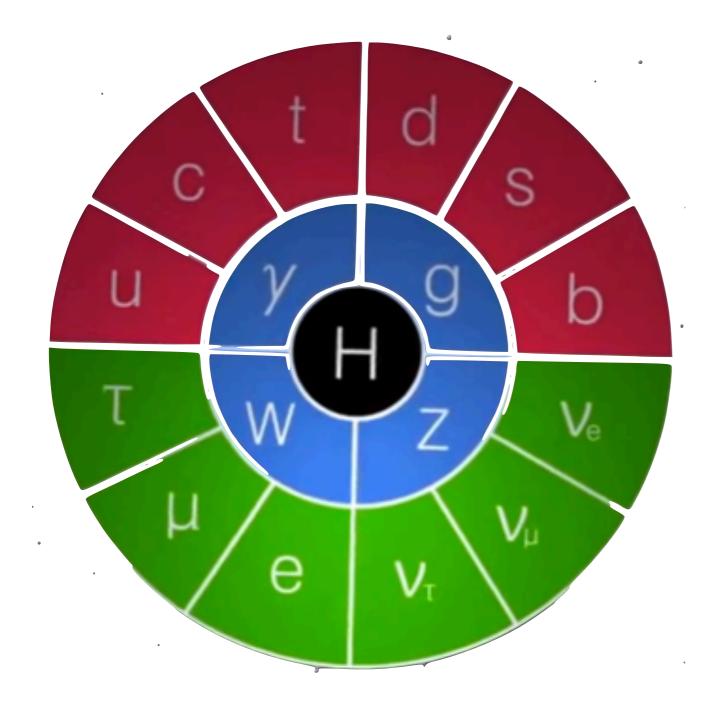


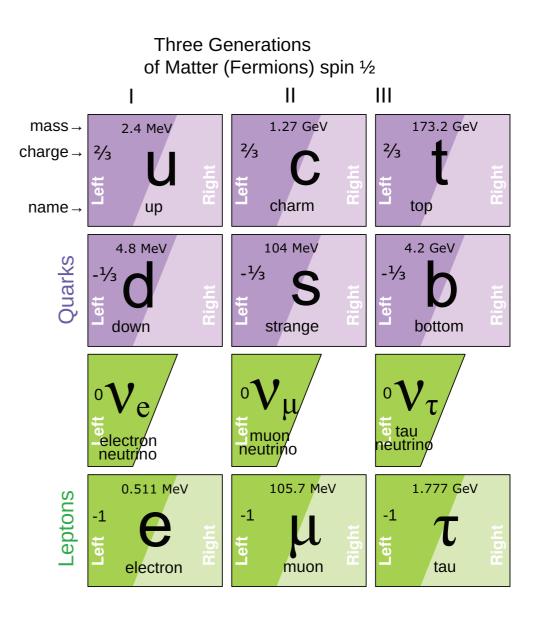
positron (e^+) discovered by C. Anderson in 1932

Christophe Grojean

The Standard Model: Matter

how many guarks and leptons?

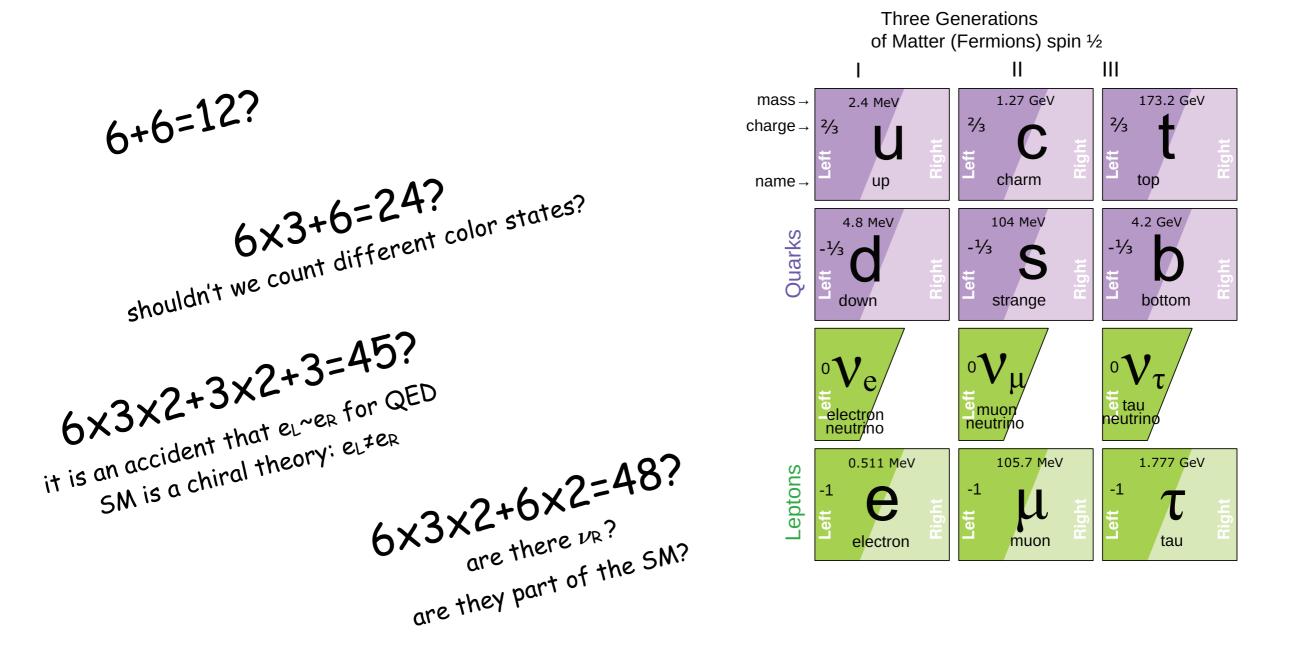




an easy question... a complicated answer!

The Standard Model: Matter

how many quarks and leptons?



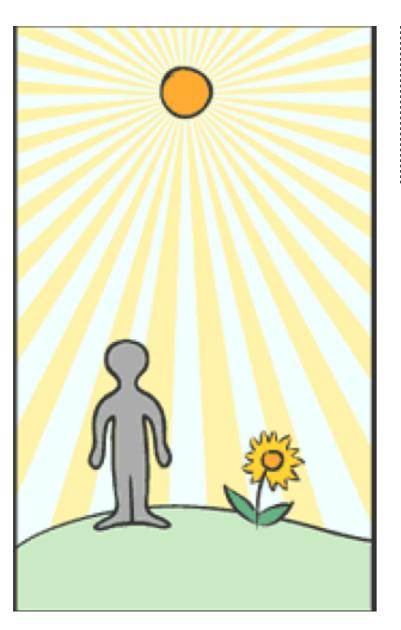
an easy question... a complicated answer!

The fundamental interactions

Interactions between Particles

(The very observation that the sun is shining for several millennia tells us that there are various mechanisms of energy production.

<u>Sun = gigantic source of energy</u>



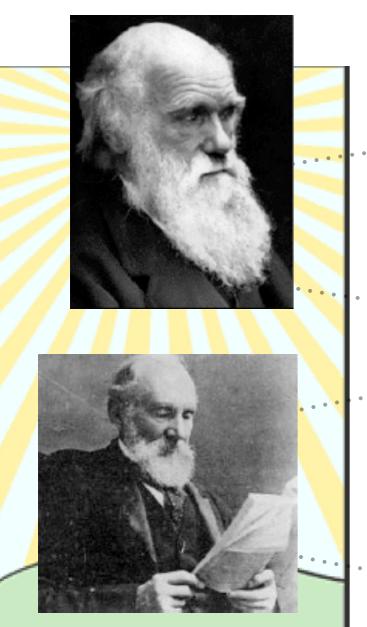
 $^{-1}$ cm³ of ice under the sun (in the summer) melts in ~ 40mn

an ice cap 1 cm thick and 300 million km of diameter centered around the sun will melt in 40mn

energy produced by burning 10¹⁹ liters of oil (~ volume Sun-Mercury/1000)

Interactions between Particles

(The very observation that the sun is shining for several millennia tells us that there are various mechanisms of energy production.



<u>Sun = gigantic source of energy</u>

Darwin ("On the origin of species by means of natural selection", 1st edition, 1859) estimates that the age of the Earth, and thus the age of the Sun also, has to be larger than 300 millions of years to account for the erosion of hills in South England.

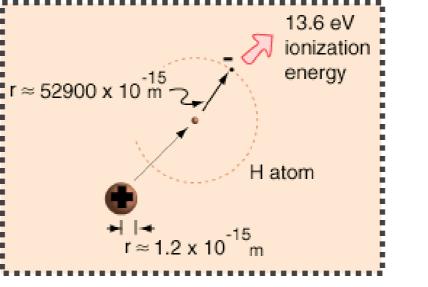
Thompson, Lord Kelvin, computes the gravitational energy of the Sun and with the assumption that it is entirely converted in heat, concludes that the Sun cannot be older than 20 million years (chemical energy would allow the Sun to shine for at most 3000 years)

15

We know today that the Sun is more than 4.5 billion years

Christophe Grojean

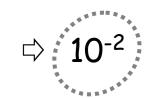
Different Interactions





mass of an atom = mass of nucleus + masses of electrons <u>example</u> : hydrogen atom, mass ~ 1 GeV, binding energy ~ 13 eV
• 10⁻⁸
• Muclear Physics

mass of a nucleus < Σ masses of protons and neutrons <u>example</u> : Helium nucleus, mass ~ 4 GeV, binding energy ~ 28 MeV



CERN, July 2015

28,300,000 eV

to break apart

α

protons and neutrons.

Particle Physics

mass of a proton or a neutron $\gg \Sigma$ masses of quarks proton mass ~ 1 GeV, constituent quarks masses~ 12 MeV \Rightarrow 10²

Christophe Grojean

The Standard Model: Interactions Even though EM is way stronger than gravity, it was unnoticed until ~ 300 years because 1-1=0

electromagnetic interactions

(1873, Maxwell)

tested with an accuracy of 10⁻⁸

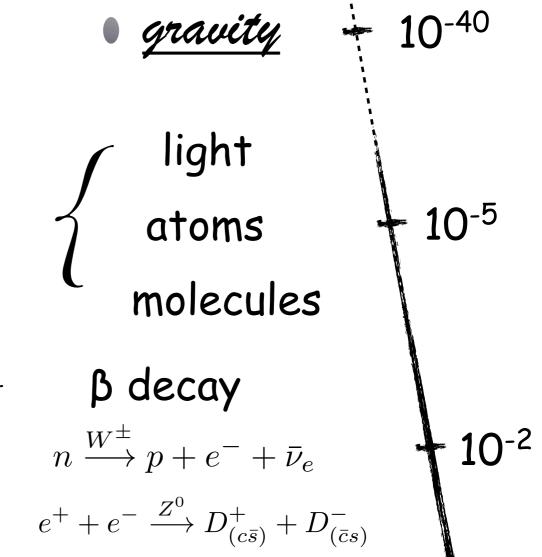
weak interactions

(1933, Fermi)

tested with an accuracy of 10⁻³

strong interactions

tested with an accuracy of 10⁻¹



atomic nuclei a decay $^{238}_{92}U \rightarrow ^{234}_{90}Th + ^{4}_{2}He$

17

(1911, Rutherford ; 1921, Chadwick and Biesler)

Christophe Grojean

BSM

strer