

# Dark Photons

GOOGLE dark matter pie chart

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Eric

SafeSearch

Today

Dark Energy	71.4%
Dark Matter	24%
Atoms	4.6%

Tomorrow

Dark Energy	75%
Dark Matter	21%
Normal Matter	4%

What The Universe Is Made Of

Dark Energy	68.3%
Dark Matter	26.8%
Atoms	4.9%

Today

Dark Energy	73%
Dark Matter	23%
Luminous Matter	3.6%
Atoms	0.4%

What The Universe Is Made Of

Dark Energy	73%
Dark Matter	23%
Other nonluminous components	0.3%
Luminous matter	0.4%
Atoms	0.004%

Stuff in the Universe

Dark Matter	23%
Dark Energy	73%
Other nonluminous components	0.3%
Luminous matter	0.4%
Atoms	0.004%

What The Universe Is Made Of

Dark Energy	72%
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Neighbors

Dark Matter	63%
Neighbors	10%

Dark Matter

Dark Matter	63%
-------------	-----

Dark Matter

Dark Matter	63%
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Dark Energy

Dark Energy	72%
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# What is a dark photon

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- Dark Matter is generally described as a WIMP.
- But it need not be that boring!
- One can consider different types of dark matter particles that interact and coexist.
- The dark photon is the general descriptor for the thingy that mediates this dark matter interaction.

# What do we know about DM?

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- It's dark
- It's cold
- It's collisionless (mostly)
- It doesn't annihilate (much)
- It's NOT necessarily non-interacting

# The EM like Dark Photon

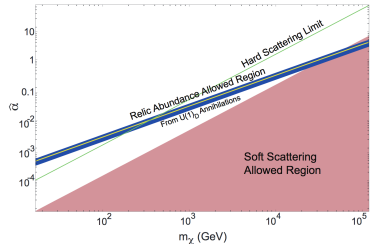
Dark Photons

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- When originally proposed in (arXiv:0810.5126) the dark photon ( $\hat{\gamma}$ ) was the force carrier for "dark charge"
- It was the gauge boson of a  $U(1)_D$  symmetry
- It was massless
- Dark matter could have +1,-1, or 0 dark charge
- The dark matter with dark photons model was characterized by  $M_X$ , the mass of a dark matter particle, and  $\hat{\alpha}$ , the dark fine structure constant.
- The dark photon had no direct SM interactions

$$SU(3)_C \times SU(2)_L \times U(1)_Y \longrightarrow \\ SU(3)_C \times SU(2)_L \times U(1)_Y \times U(1)_D$$

- Here is the allowed regions as calculated in arXiv:0810.5126 based only upon cosmological observations
- The relic abundance allowed region applies to models in which  $U(1)_D$  is the only force coupled to the dark matter; in models where the DM is also weakly interacting, this provides only an upper limit on  $\hat{\alpha}$

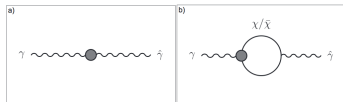


# ElectroDarkWeak

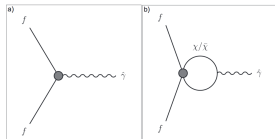
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- You can have a DM particle that interacts weakly AND through  $\hat{\gamma}$



- But if you do almost all mixing ends up cancelling out

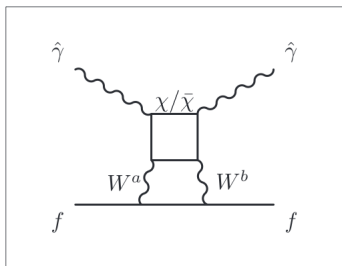


# ElectroDarkWeak

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- The lowest order SM coupling is at order  $\lambda_f \alpha^2 \hat{\alpha}$  where  $\lambda_f$  is the coupling of the fermion involved
- In this scenario a signal in a direct detection experiment would be identical to that of normal (non-interacting) dark matter.





Quoth arXiv:0810.5126:

*Therefore, the presence of a new unbroken  $U(1)_D$  in the dark sector could only be probed via its effect on galactic dynamics”*

Moving on to the (more recent) paper du jour

## Constraints on Dark Photon from Neutrino-Electron Scattering Experiments

- As the title suggests the authors consider neutrino elastic scattering as the main channel by which dark photons can be observed.
- But wait. Didn't we just hear that terrestrial detection type experiments are hopeless?
- It seems since the original proposition of the dark photon idea has evolved and expanded.

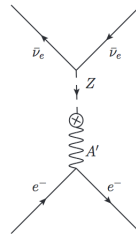
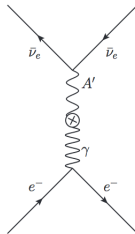
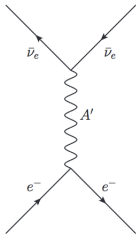
# The dark photon according to arxiv:1502.07763

- $\hat{\gamma}$  can have non-zero but tiny mass.
- There are two ideas through which the dark photon can be considered.
  - 1 The dark photon comes from a  $U(1)'$  symmetry that mixes with the SM  $U(1)_Y$  symmetry
    - This results in a kinetic mixing between the the photon and the dark photon
    - This has the form  $\epsilon B'_{\mu\nu} F^{\mu\nu}$  where  $\epsilon$  is the strength of the mixing
  - 2 The other option for getting the dark photon is through a  $U(1)_{B-L}$  symmetry
    - This paper chooses to consider the  $U(1)_{B-L}$  primarily.
    - The  $U(1)_{B-L}$  model is characterized by dark photon mass  $M_{A'}$  and the coupling constant  $g_{B-L}$

# $\hat{\gamma}$ Interactions

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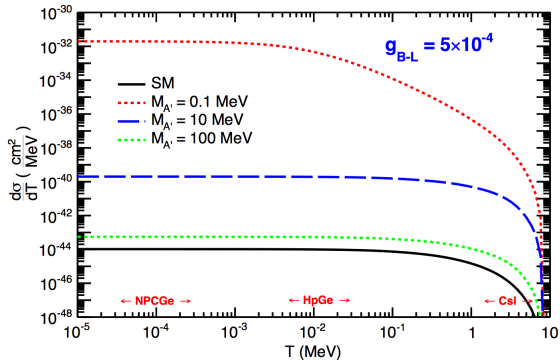


# Elastic Scattering Spectrum

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Since you now have all these extra diagrams for neutrino-electron elastic scattering you expect the cross section to change.

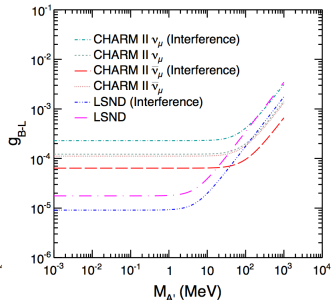
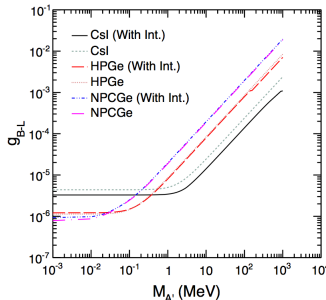


# Interference

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- The authors of the paper harped on about the importance of including interference effects
- Apparently there were some people before who hadn't done this



The authors made the world's 2nd most confusing plot to illustrate the error of this neglect

# A quick summary of experiments

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Several experiments were considered

- TEXONO - Reactor Anti-Neutrinos
- LSND - Accelerator Neutrinos
- BOREXINO - Solar Neutrinos
- GEMMA - Reactor Anti-Neutrinos
- CHARM II - Accelerator Muon Neutrinos

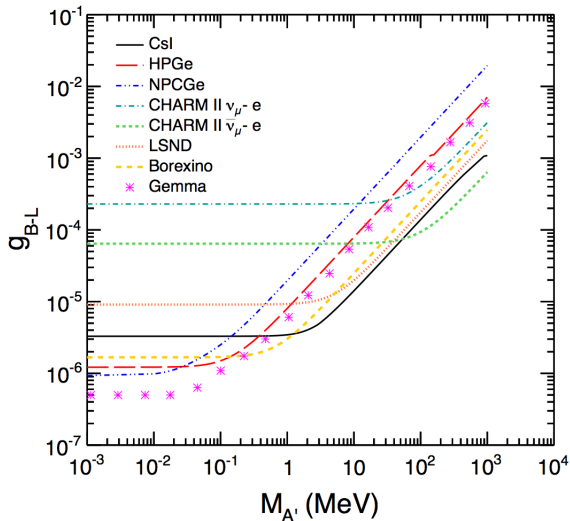
TEXONO comes in 3 different flavors, CsI, HPGe, NPCGe which denote different detector arrays that were swapped in or out. Each different detector had differing physics advantages.



# Exclusion Plot

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# Global Dark Photon Exclusion

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