Neutrinos and Nucleosynthesis in Gamma Ray Bursts (black hole accretion disks)

NIC 2006

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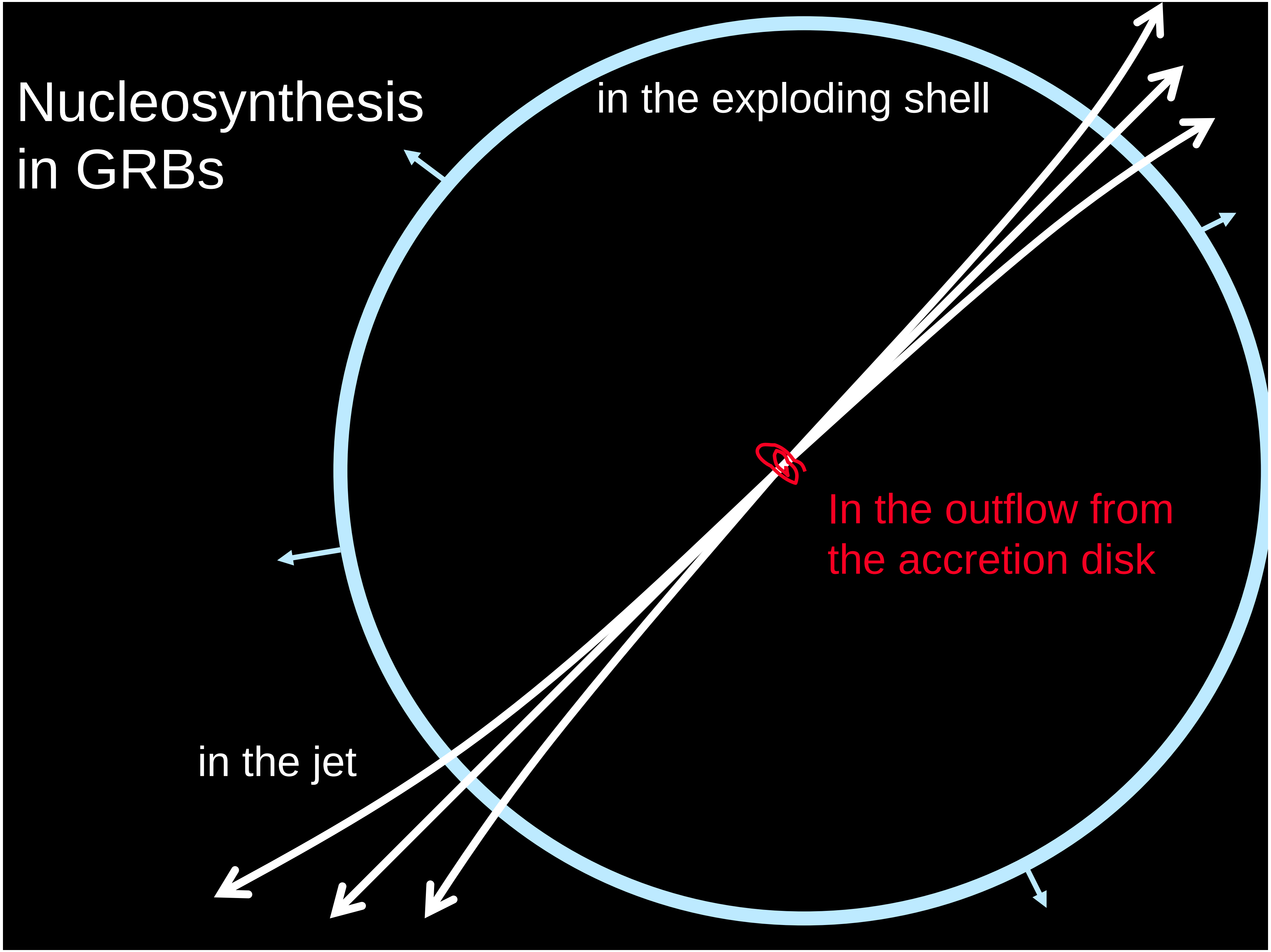
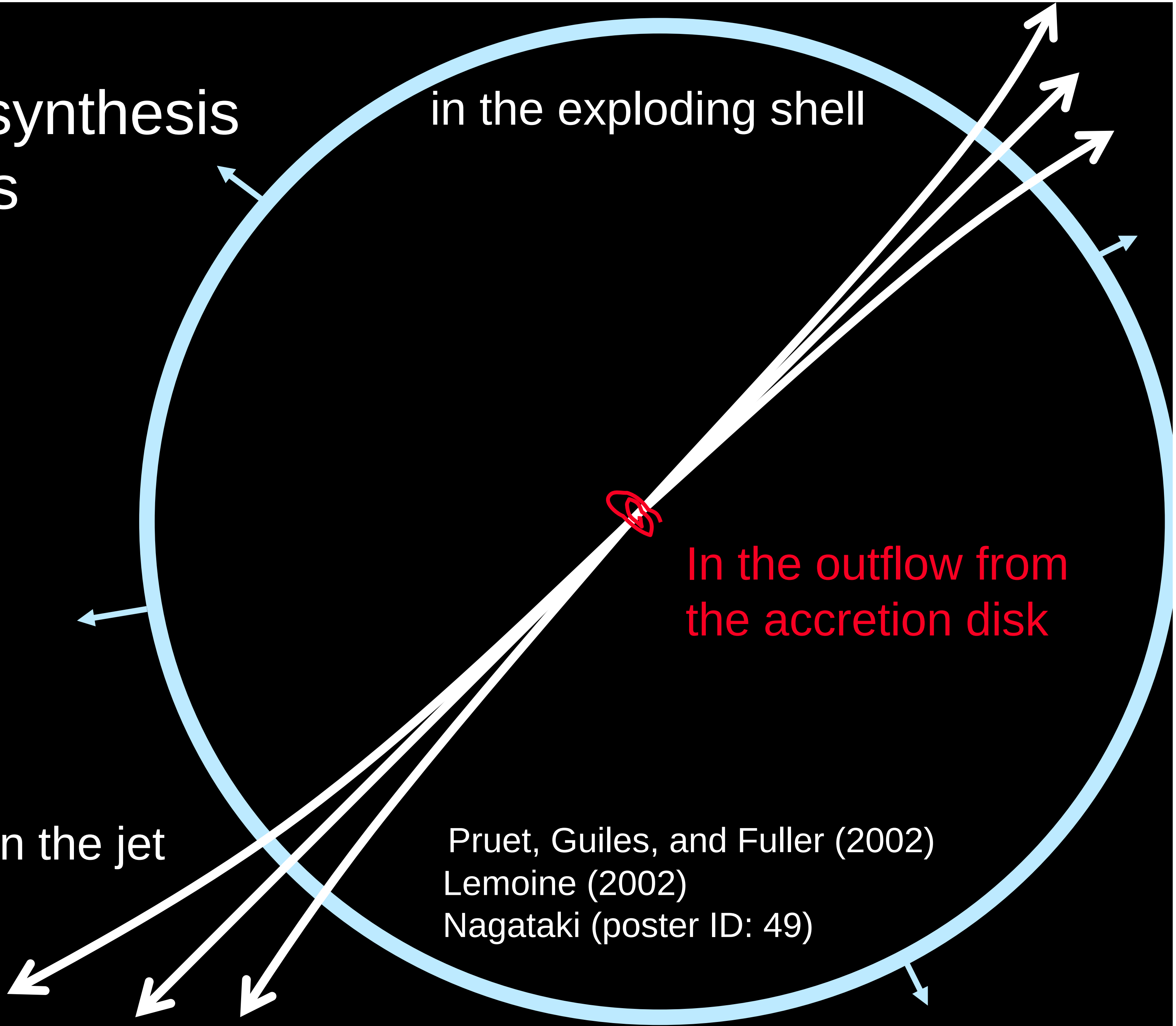
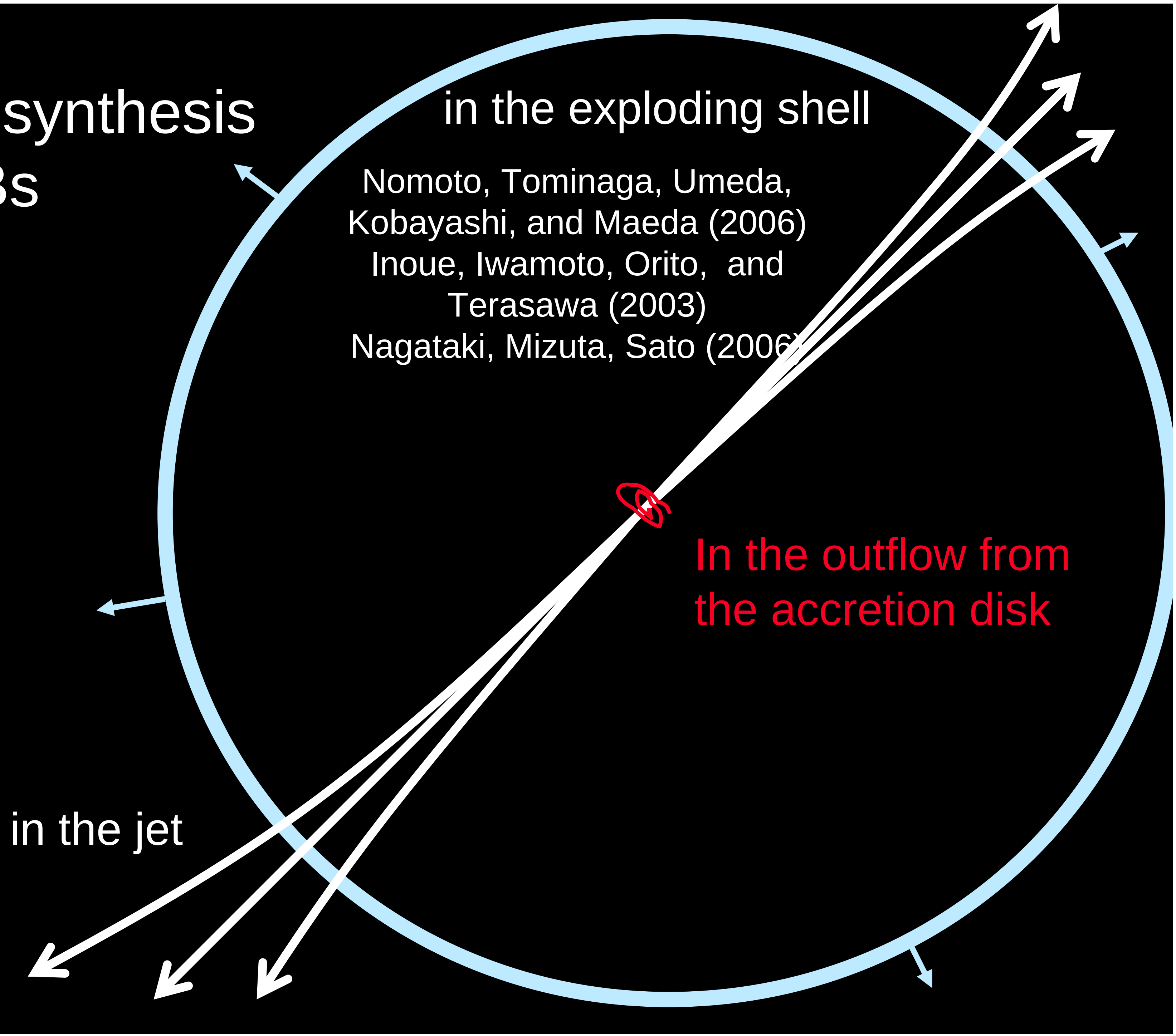

Nucleosynthesis
in GRBs
in the exploding shell
In the outflow from
the accretion disk
in the jet

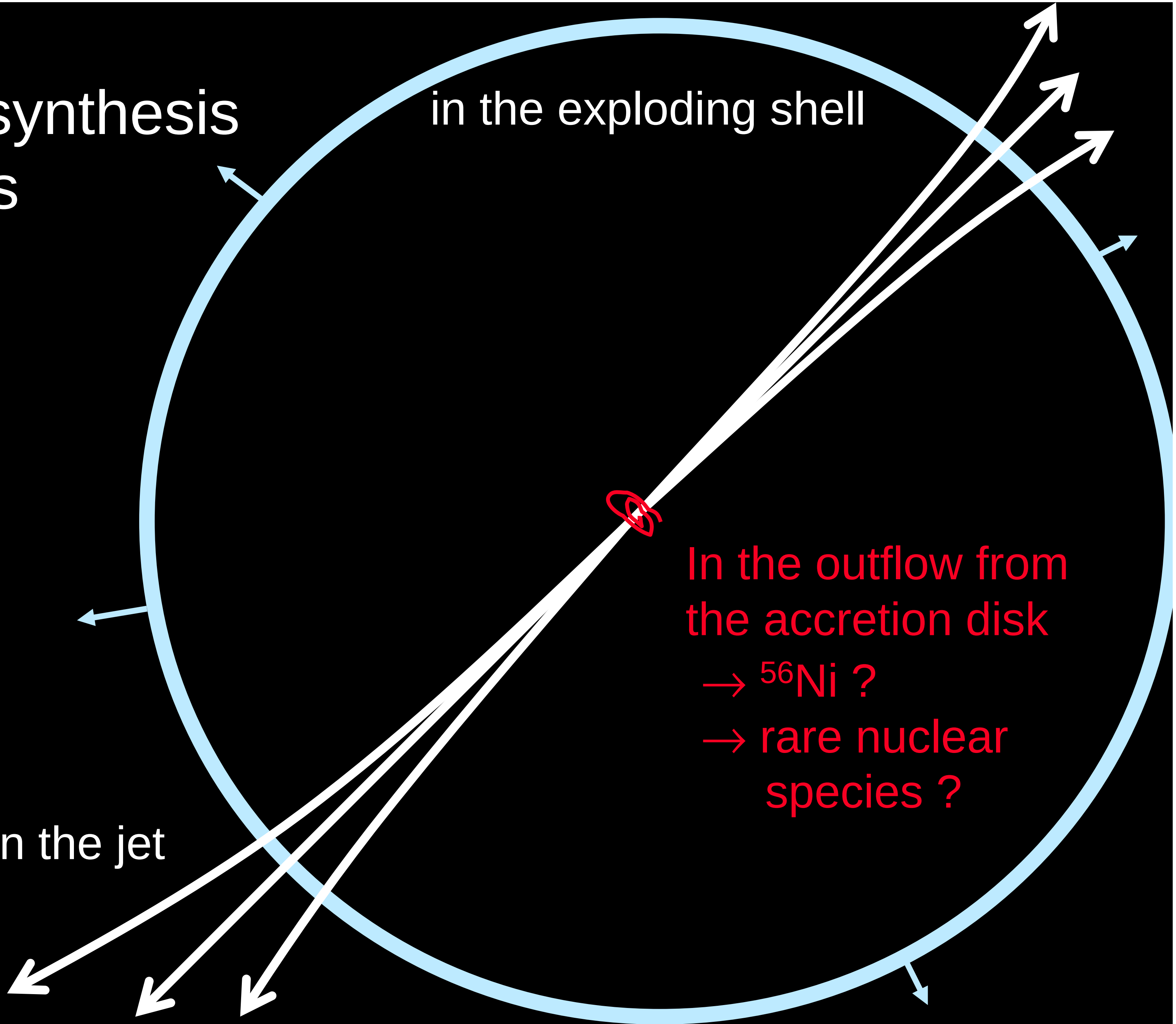
Nucleosynthesis in GRBs



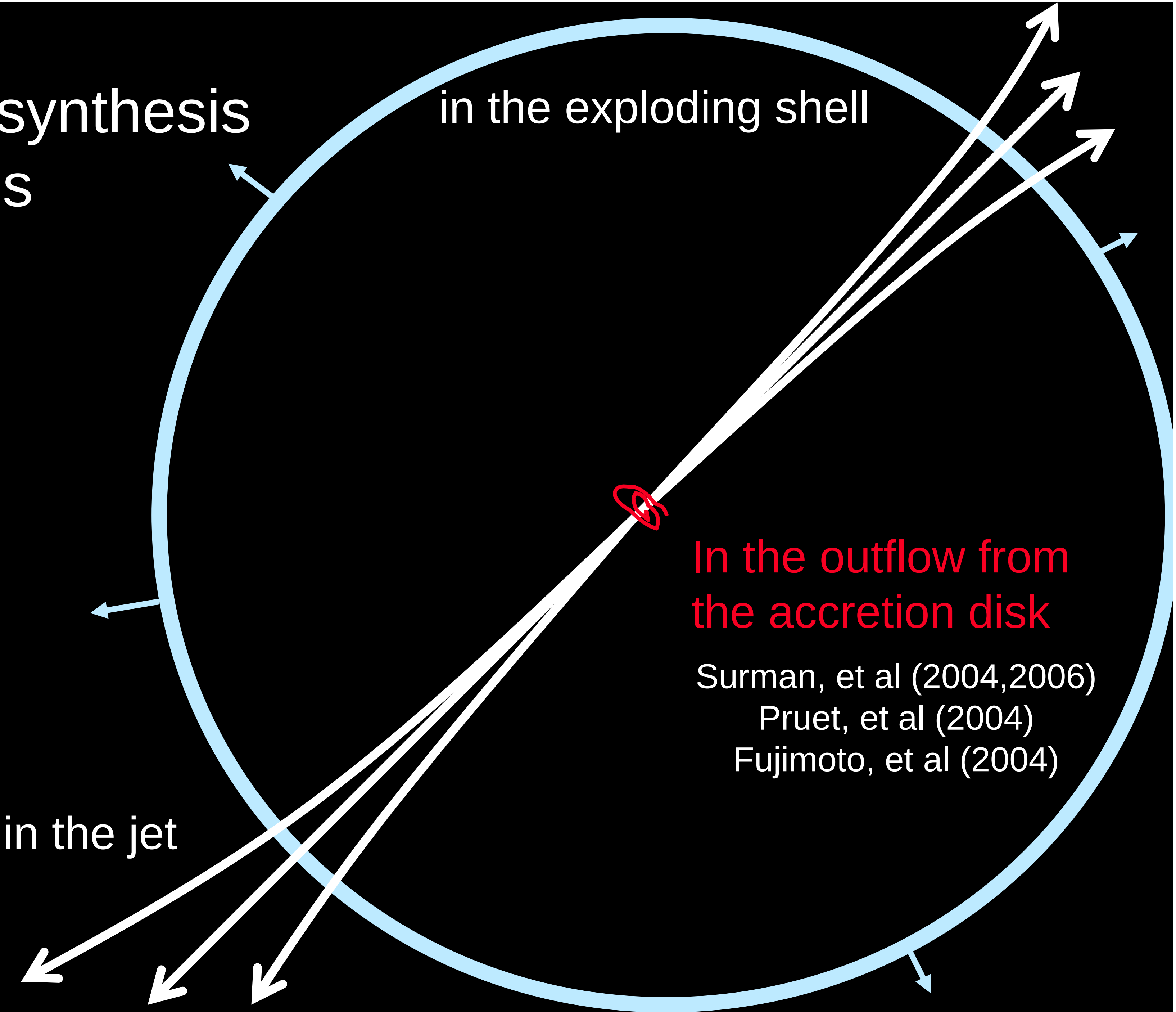
Nucleosynthesis in GRBs



Nucleosynthesis in GRBs

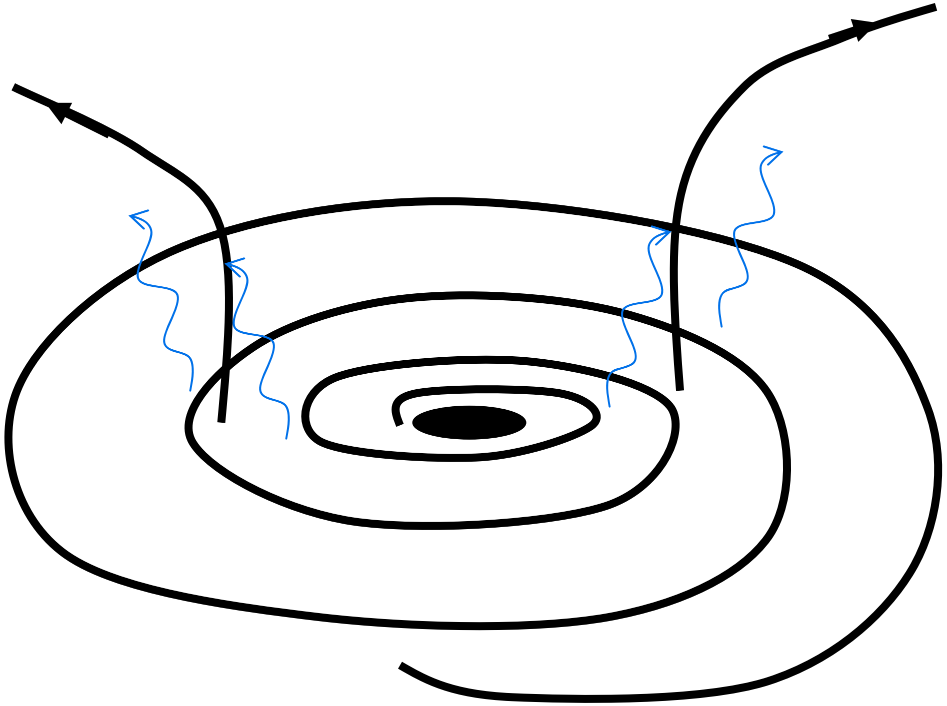


Nucleosynthesis in GRBs



Nucleosynthesis in the Outflow

Follow material
- through disk
- as ejected in outflow



Electron fraction set by:

$$e^- + p \leftrightarrow n + \nu_e$$

$$e^+ + n \leftrightarrow p + \overline{\nu}_e$$

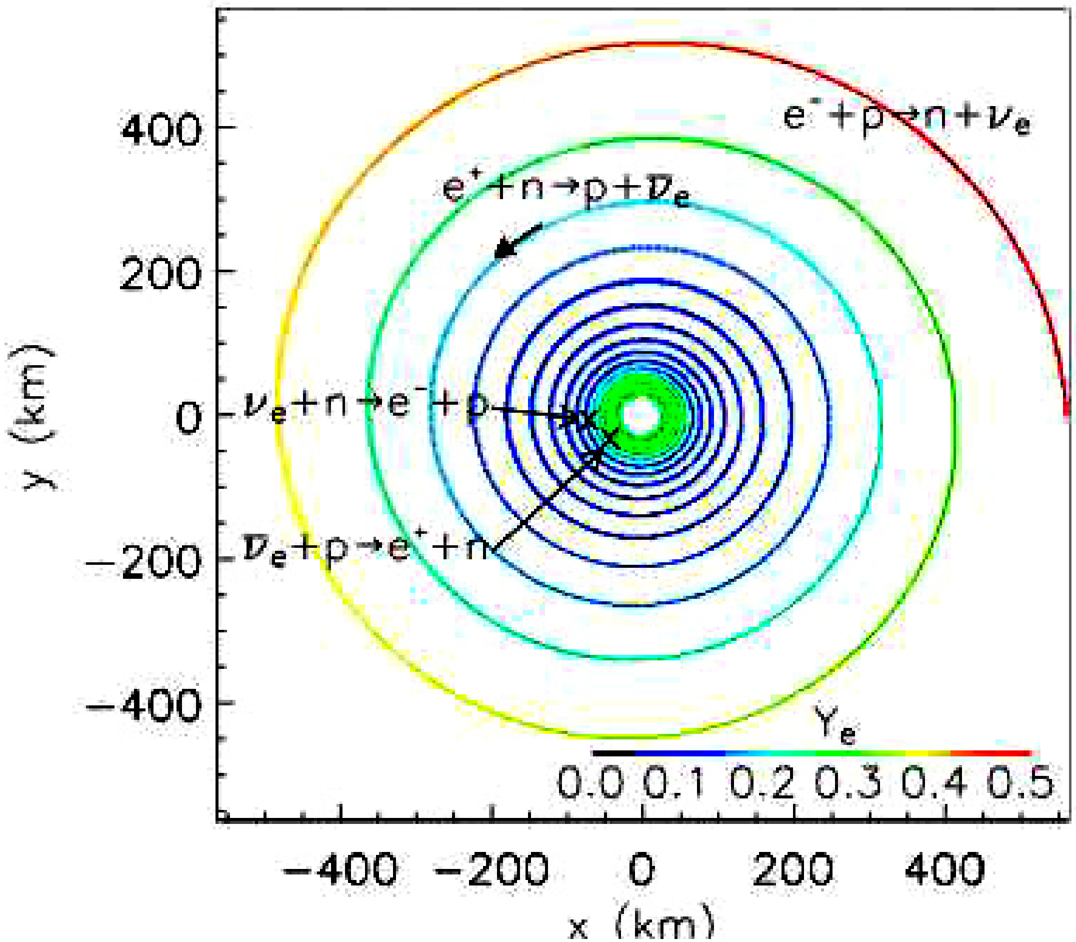
Disk models:

Low accretion rate disks - $\dot{m} < 1$, where $\dot{m} = 1 \Rightarrow$ 1 solar mass/second

Popham, Woosley, and Fryer (1999)

High accretion rate disks - $\dot{m} \geq 1$

DiMatteo, Perna, and Narayan (2002)



Surman & McLaughlin, ApJ, 603, 611 (2004)

Evolution of Y_e
in the disk

DPN $\dot{m} = 1.0$

Outflow Parameterization

Take velocity as a function of radial distance from the black hole to be

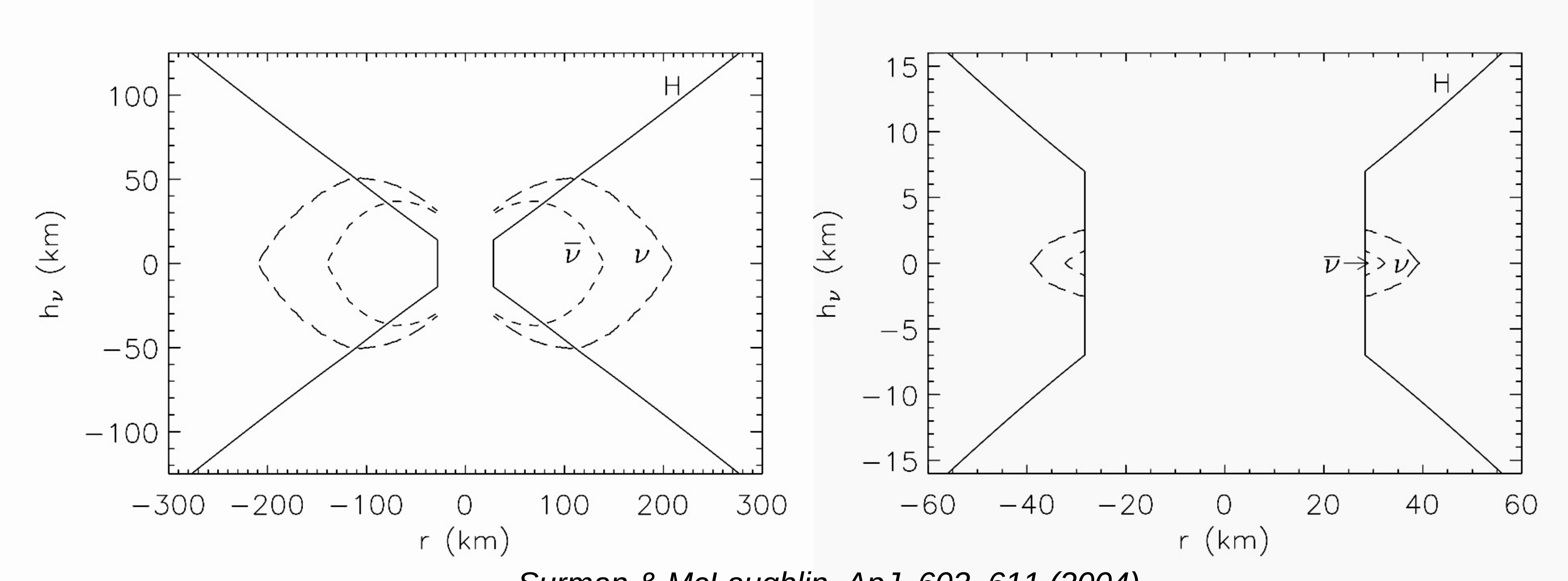
$$u = v_{\infty}\left(1 - \frac{R_o}{R}\right)^{\beta}$$

where 5,000 $< v_{\infty} <$ 50,000 km/s, 0.2 $< \beta < 3.0$

Take flow to be vertical at first, then radial

Consider adiabatic flows with entropy 10 $< s <$ 50

Neutrino Decoupling Surfaces



Surman & McLaughlin, ApJ, 603, 611 (2004)

Neutrino flux coming from disk is dominated by contribution from optically thick region

⇒ When antineutrino surface is large, the antineutrinos tend to dominate (higher T_ν)

$$\bar{\nu}_e + p \rightarrow n + e^+$$

⇒ If antineutrino surface is small or nonexistent, neutrino flux dominates

$$\nu_e + n \rightarrow p + e^-$$

Nuclear Recombination in the Outflow

Full nuclear network code:

W. R. Hix, *J. Comp. Appl. Math*., 109, 321 (1999)
(J. Beun, R. Surman)

r-process nucleosynthesis code:

J. Walsh, B.S. Meyer, R. Surman

Look for: * ^{56}Ni

* rare nuclear species:

p process

r process

Low accretion rate => Nickel Synthesis

PWF

$\dot{m} = 0.1$

$r_o = 100 \text{ km}$

$v_\infty = 0.1c$

QuickTime™ and a
TIFF (Uncompressed) decompressor
are needed to see this picture.

0.5 0.4 0.3 0.2 0.1

^{56}Ni mass fraction

Lines indicate Ye

Surman, McLaughlin & Hix, ApJ, 643, 1057 (2006)

Moderate accretion rates => Nickel Synthesis

DPN

$\dot{m} = 1.0$

$r_o = 250\ \text{km}$

$v_\infty = 0.1c$

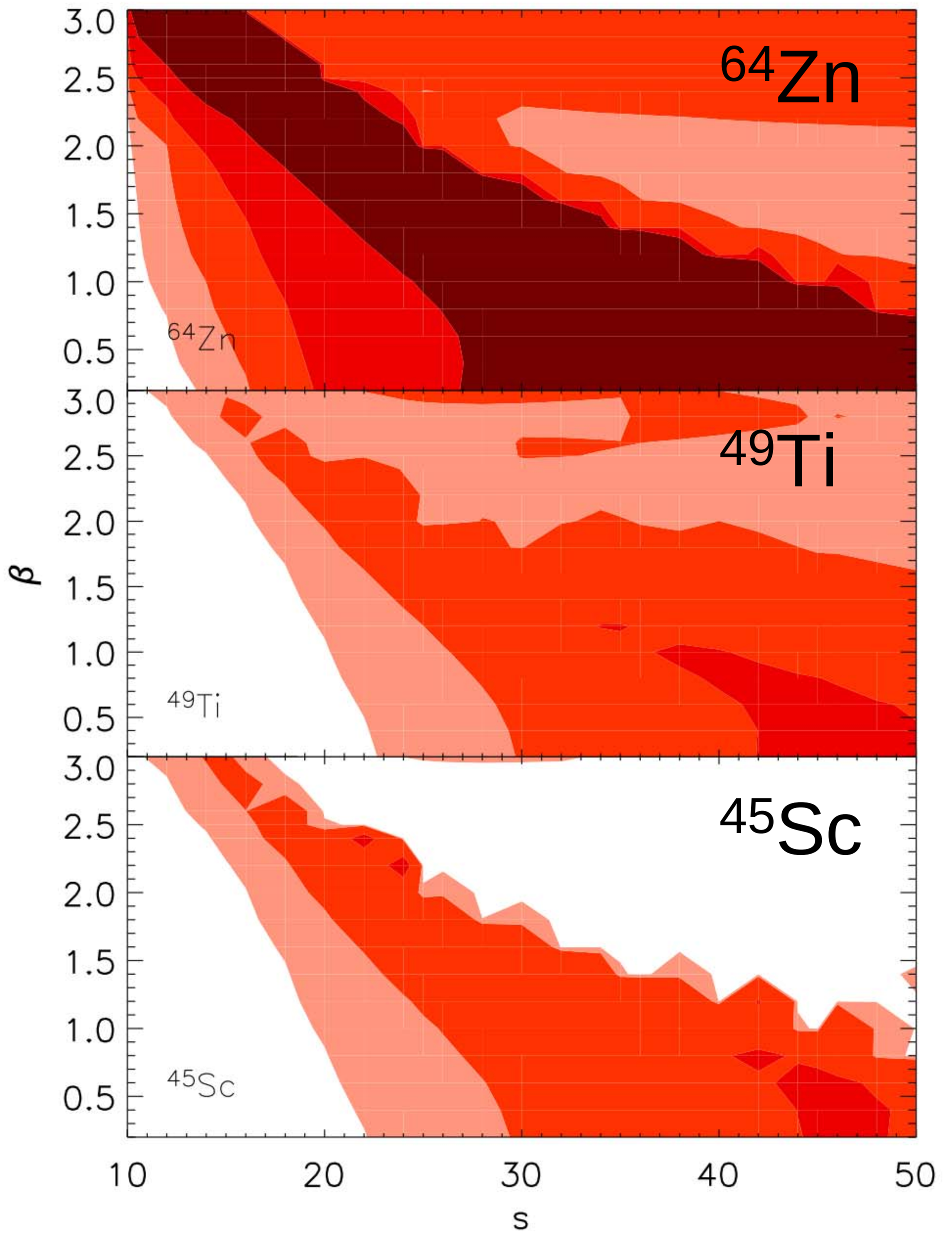
QuickTime™ and a
TIFF (Uncompressed) decompressor
are needed to see this picture.

0.5 0.4 0.3 0.2 0.1

^{56}Ni mass fraction

Surman, McLaughlin & Hix, ApJ, 643, 1057 (2006)

Overproduction Factors



Surman, McLaughlin, & Hix, ApJ, 643, 1057 (2006)

$$O(j) = \left(\frac{M_{wind}}{M_{SN\ ejecta}} \right) \times \left(\frac{X_{wind}}{X_{solar}} \right)$$

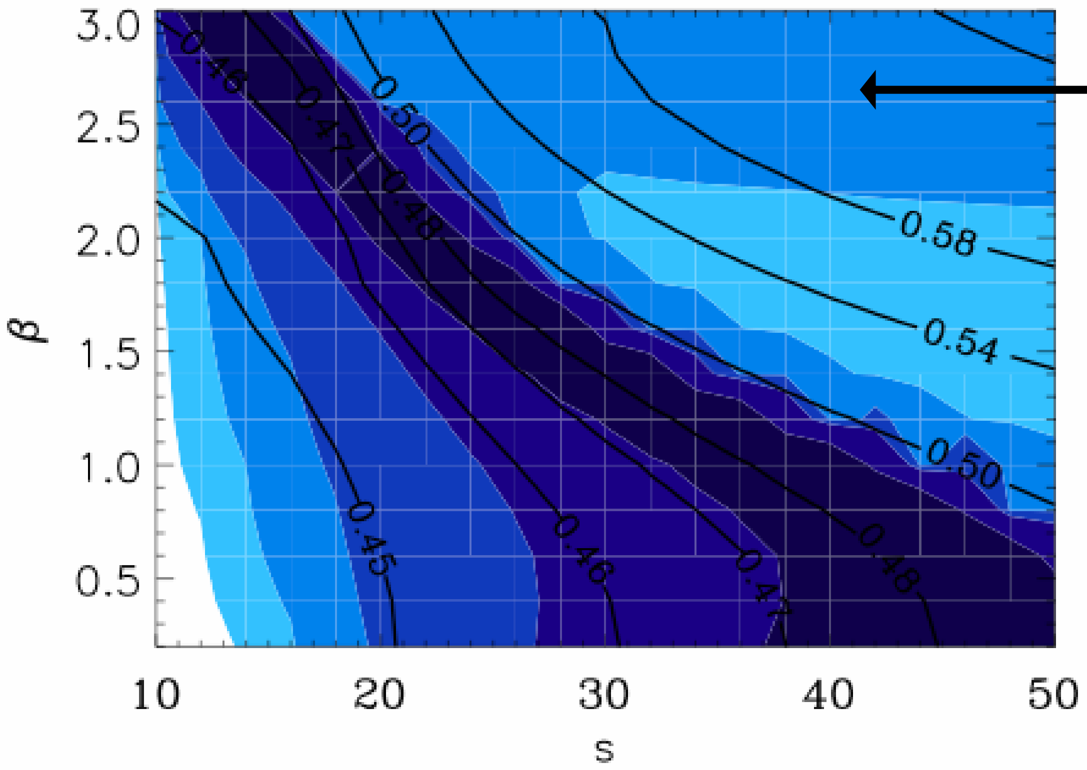
PWF

$\dot{m} = 0.1$

$r_o = 100$ km

$v_\infty = 0.1c$

^{64}Zn



Slow outflow - enhancement due to neutrino interactions

PWF

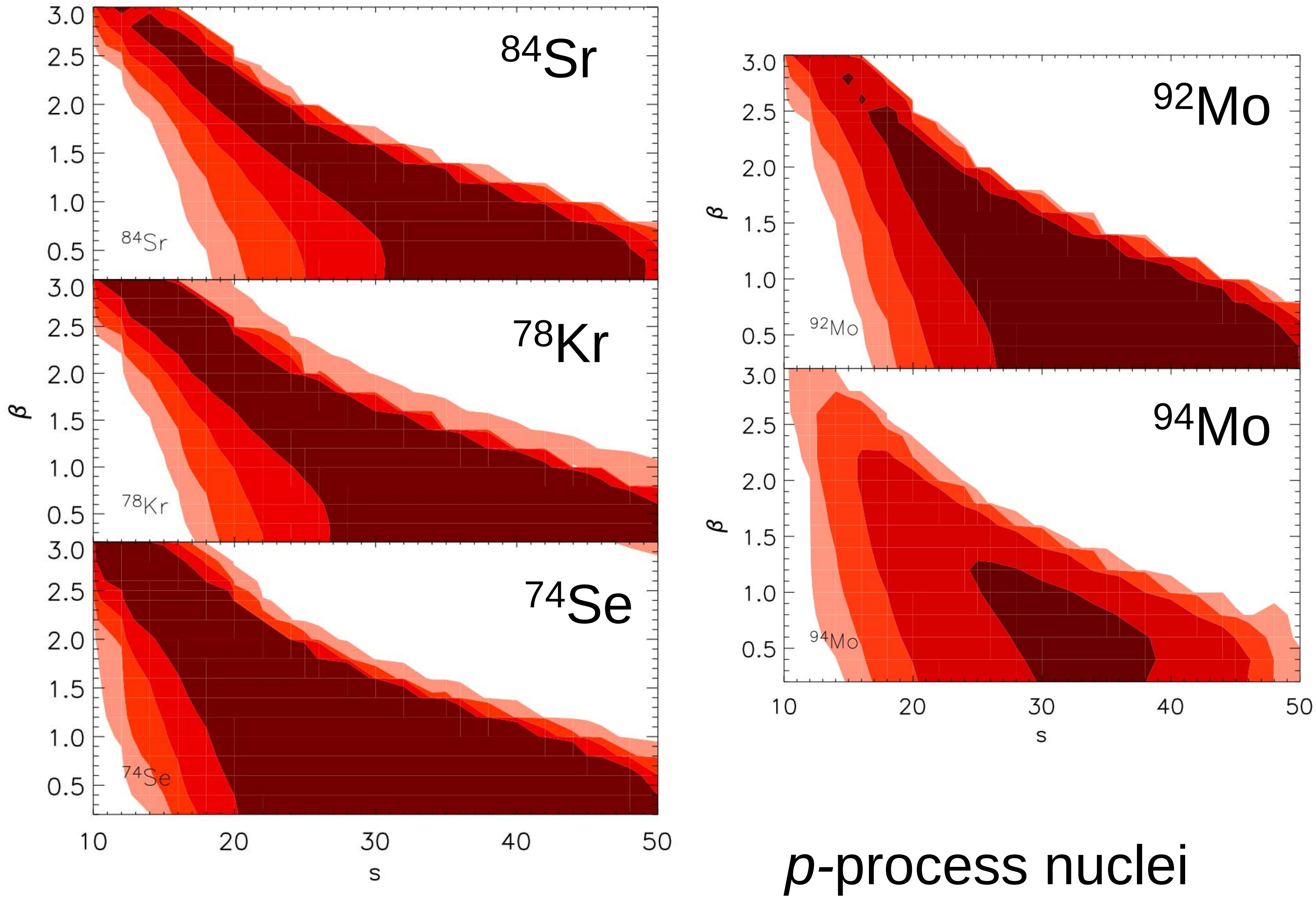
$\dot{m} = 0.1$

$r_o = 100$ km

$v_\infty = 0.1c$

O(j) >

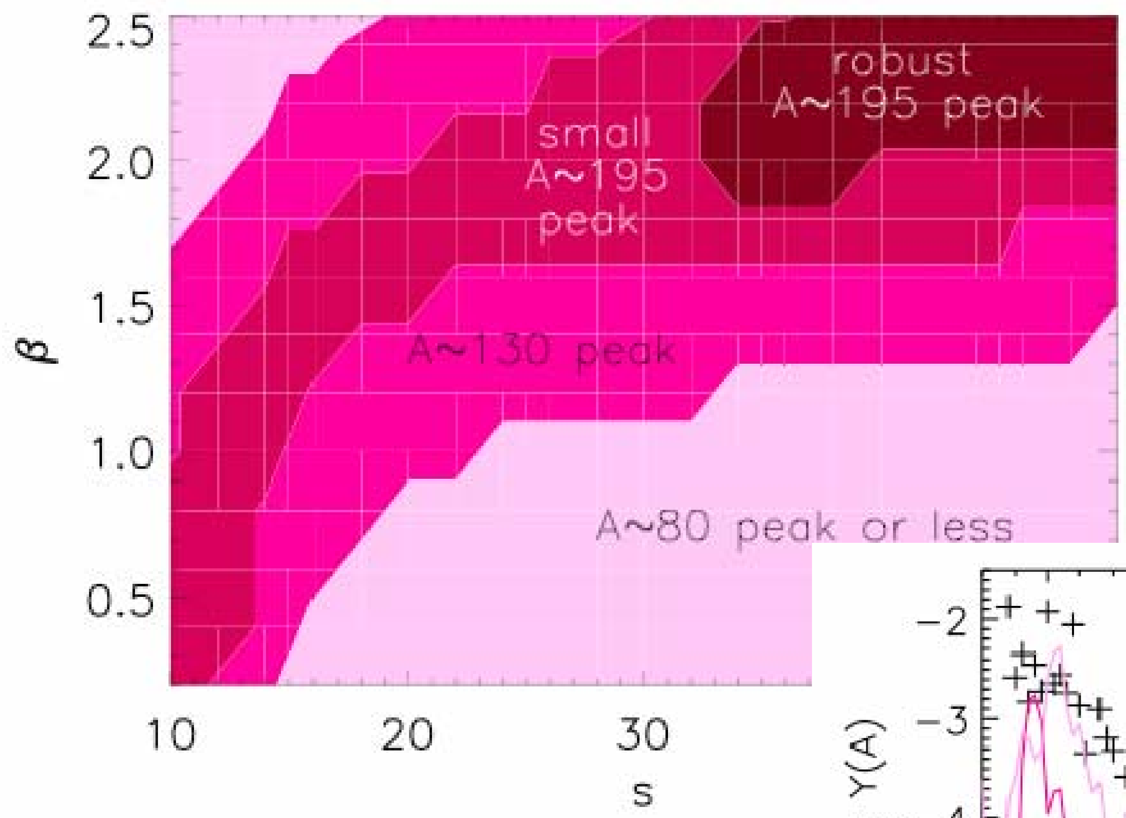
5000 1000 100 10 1



Surman, McLaughlin, & Hix, ApJ, 643, 1057 (2006)

p-process nuclei

High accretion rate => *r*-Process Nucleosynthesis



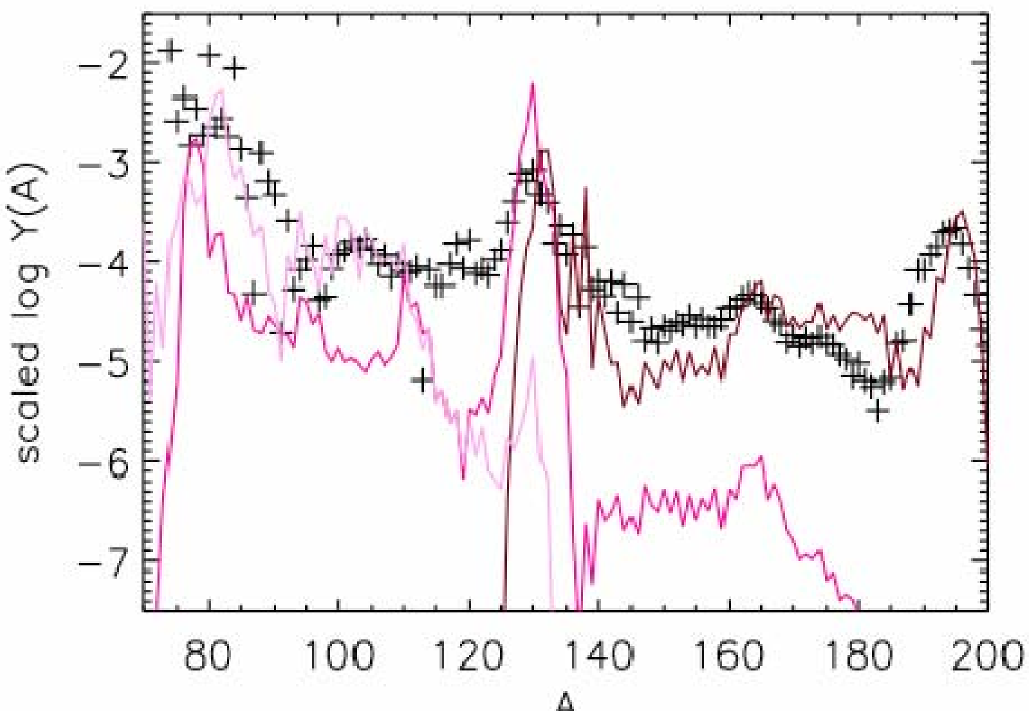
DPN

$\dot{m} = 10$

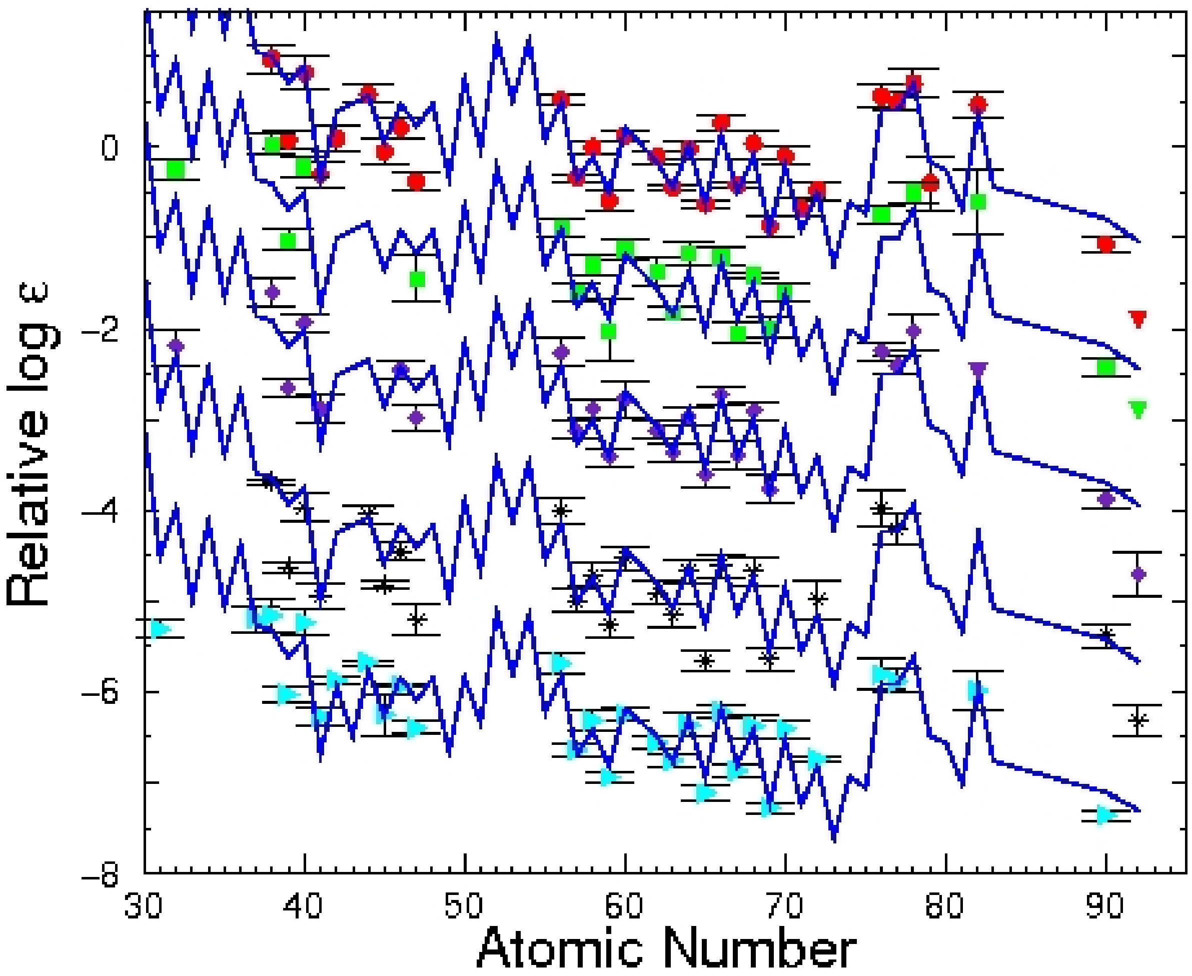
$r_o = 250 \text{ km}$

$v_\infty = 0.1c$

Surman, McLaughlin, and Hix, ApJ, 643, 1057 (2006)



Metal-poor Halo Star data



From J. Cowan's talk, 6/27/06

Conclusions

Given disk and outflow parameters, we can determine what nucleosynthesis will result from an understanding of the neutrinos

Nucleosynthesis in accretion disk outflows provides a promising mechanism for GRB nickel production

Additionally, GRBs may contribute to the galactic abundances of certain rare nuclear species, such as *r*-process nuclei or light/intermediate mass proton-rich nuclei such as ^{64}Zn, ^{45}Sc, and ^{92}Mo