Large scale structure of the Universe: the angular power spectrum and bispectrum

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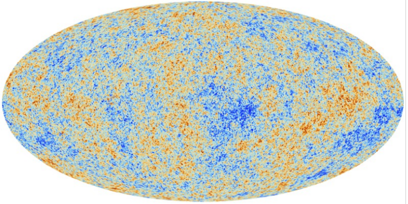
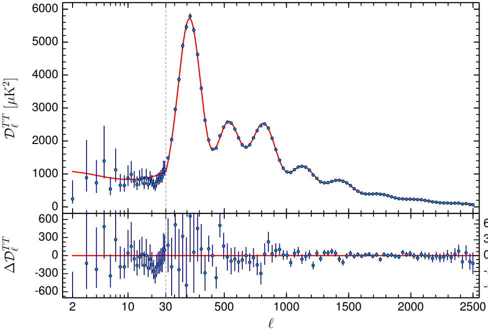
WARSAW, POLAND

Outline

The CMB



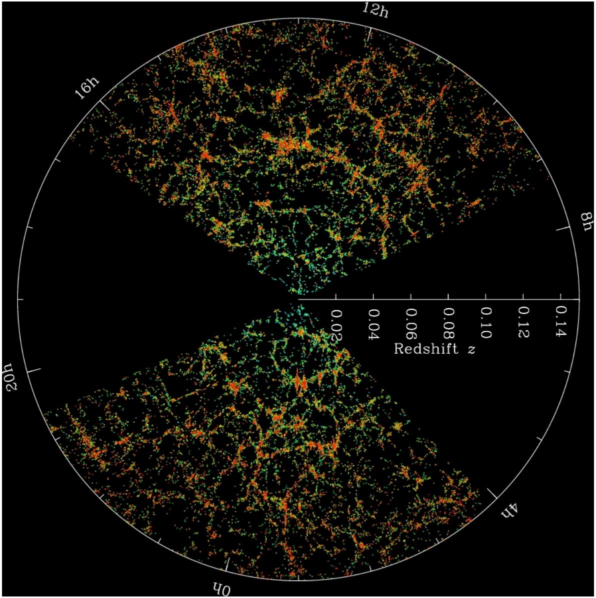
CMB sky as seen by Planck



$D_\ell = \ell(\ell+1)C_\ell/(2\pi)$

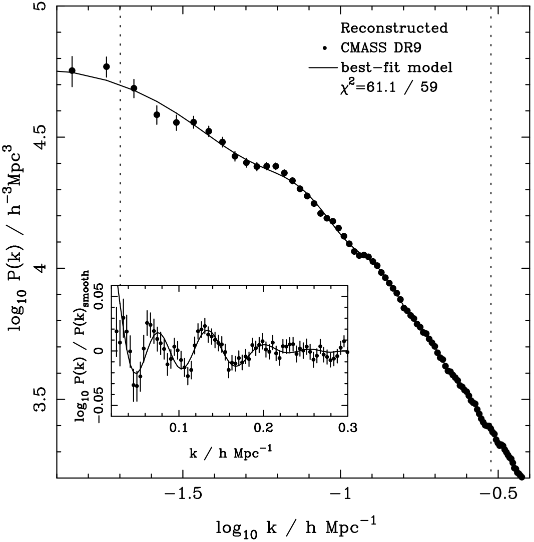
The Planck Collaboration: Planck results 2015 XIII

Introduction



M. Blanton and the Sloan Digital Sky Survey Team.

Galaxy power spectrum from the Sloan Digital Sky Survey (BOSS)



from Anderson et al. ’12

SDSS-III (BOSS) power spectrum.

Galaxy surveys $\simeq$ matter density fluctuations, biasing and redshift space distortions.

But...

- We have to take fully into account that all observations are made on our past lightcone which is itself perturbed.
  We see density fluctuations which are further away from us, further in the past.
  We cannot observe 3 spatial dimensions but 2 spatial and 1 lightlike, more precisely we measure 2 angles and a redshift.

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- But of course much more for future surveys like DES, Euclid, WFIRST and SKA.

Cosmological distances

In a Friedmann Universe the (comoving) radial distance is

$$r(z) = \int_0^z \frac{dz'}{H(z')} = \frac{1}{H_0}\int_0^z \frac{dz'}{\sqrt{\Omega_m(1+z')^3 + \Omega_K(1+z')^2 + \Omega_\Lambda}}$$

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Depending on the observational situation we measure directly $r(z)$ or

$$\begin{aligned} d_A(z) &= \frac{1}{(1+z)}\chi_K(r(z)) \quad \text{the angular diameter distance} \\ d_L(z) &= (1+z)\chi_K(r(z)) \qquad\qquad \text{the luminosity distance.} \end{aligned}$$

At small redshift all distances are $d(z) = z/H_0 + \mathcal{O}(z^2)$, for $z \ll 1$. At larger redshifts, the distance depends strongly on $\Omega_K,\ \Omega_\Lambda, \cdots$.

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- Whenever we convert a measured redshift and angle into a length scale, we make assumptions about the underlying cosmology.

What are very large scale galaxy catalogs really measuring?

We now consider fluctuations in the matter distribution and in the geometry first to linear order. (See C. Bonvin & RD [arXiv:1105.5080]; Challinor & Lewis, [arXiv:1105:5092], J. Yoo et al. 2009; J. Yoo 2010)

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$$\Delta(\mathbf{n}, z) = \frac{N(\mathbf{n}, z) - \bar{N}(z)}{\bar{N}(z)}.$$

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$$\Delta(\mathbf{n}, z) = \frac{N(\mathbf{n}, z) - \bar{N}(z)}{\bar{N}(z)} .$$

$$\xi(\theta, z, z') = \langle \Delta(\mathbf{n}, z)\Delta(\mathbf{n}', z') \rangle \,, \qquad \mathbf{n} \cdot \mathbf{n}' = \cos\theta \,.$$

This quantity is directly measurable $\Rightarrow$ gauge invariant.

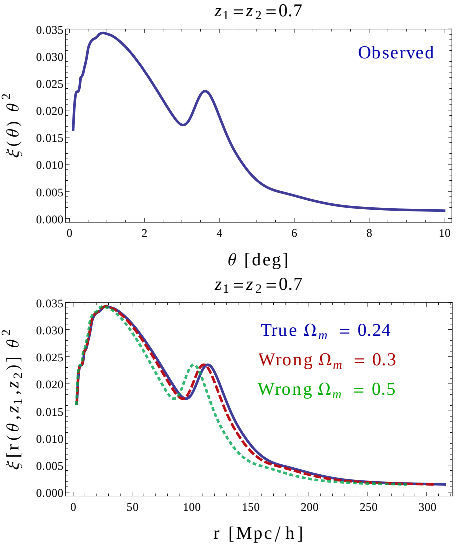
What are very large scale galaxy catalogs really measuring?

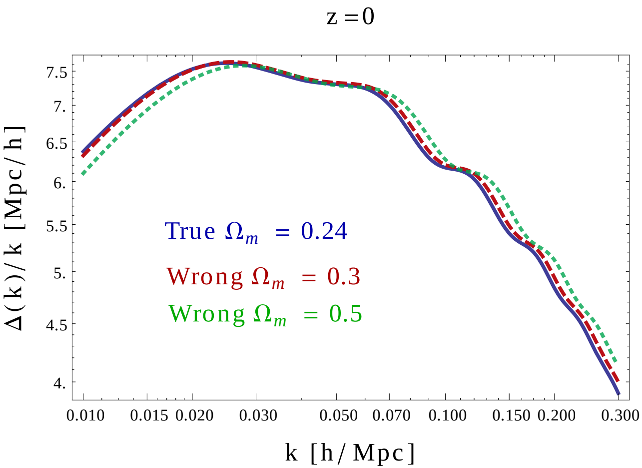
If we convert the measured $\xi(\theta, z_1, z_2)$ to a power spectrum, we have to introduce a cosmology, to convert angles and redshifts into length scales.

$$r(z_1, z_2, \theta) \stackrel{(K=0)}{=} \sqrt{r_1^2 + r_2^2 - 2r_1 r_2 \cos\theta}.$$

$$r_i = r(z_i) = \int_0^{z_i} \frac{dz}{H(z)}$$

(Figure by F. Montanari)





(Figure by F. Montanari)

$$\Delta(k)/k = k^2 P(k)$$

Putting the density and volume fluctuations together one obtains the galaxy number density fluctuations to 1st order

$$
\begin{aligned}
\Delta(\mathbf{n}, z) \quad &= \quad D_s - 2\Phi + \Psi + \frac{1}{\mathcal{H}}\left[\dot{\Phi} + \partial_r(\mathbf{V}\cdot\mathbf{n})\right] \\
&+ \left(\frac{\dot{\mathcal{H}}}{\mathcal{H}^2} + \frac{2}{r(z)\mathcal{H}}\right)\left(\Psi + \mathbf{V}\cdot\mathbf{n} + \int_0^{r(z)} dr(\dot{\Phi} + \dot{\Psi})\right) \\
&+ \frac{1}{r(z)}\int_0^{r(z)} dr\left[2 - \frac{r(z) - r}{r}\Delta_\Omega\right](\Phi + \Psi).
\end{aligned}
$$

(C. Bonvin & RD '11)

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& + \frac{1}{r(z)}\int_0^{r(z)} dr\left[2(\Phi + \Psi) - \boxed{\frac{r(z)-r}{r}\Delta_\Omega(\Phi + \Psi)}\right].
\end{aligned}
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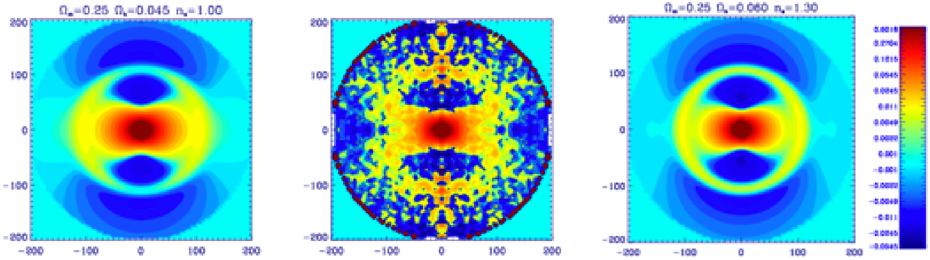
(C. Bonvin & RD '11)

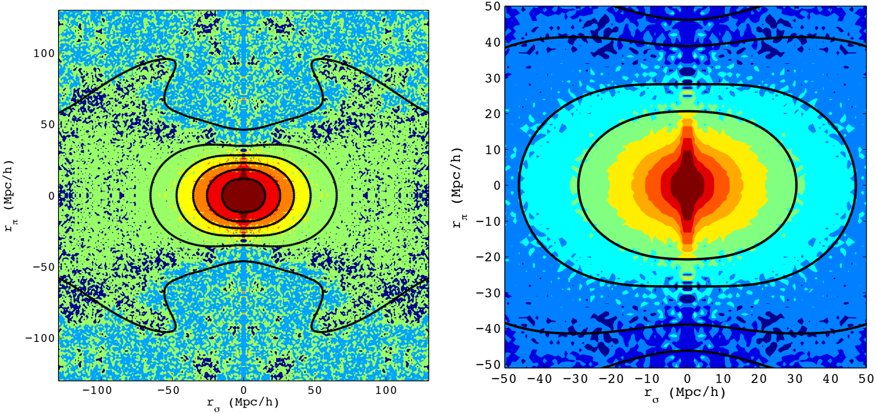


(From Gaztanaga et al. 2008)

The correlation function is not isotropic ⇒ redshift space distortions.

Redshift space distortions in the BOSS survey

(from Reid et al. '12)



The angular power spectrum of galaxy density fluctuations

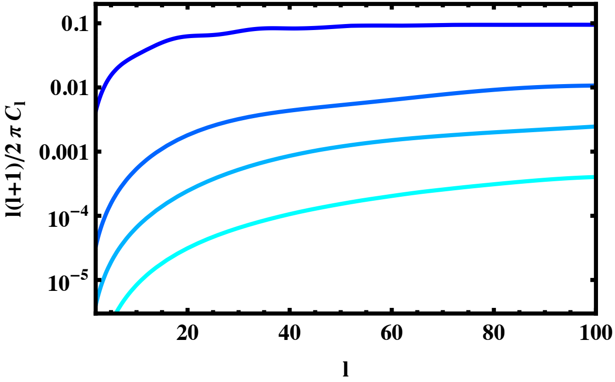
For fixed z, we can expand $\Delta(\mathbf{n}, z)$ in spherical harmonics,

$$\Delta(\mathbf{n}, z) = \sum_{\ell m} a_{\ell m}(z) Y_{\ell m}(\mathbf{n}), \qquad C_\ell(z, z') = \langle a_{\ell m}(z) a^*_{\ell m}(z') \rangle.$$

$$\xi(\theta, z, z') = \langle \Delta(\mathbf{n}, z) \Delta(\mathbf{n}', z') \rangle = \frac{1}{4\pi} \sum_\ell (2\ell + 1) C_\ell(z, z') P_\ell(\cos\theta)$$

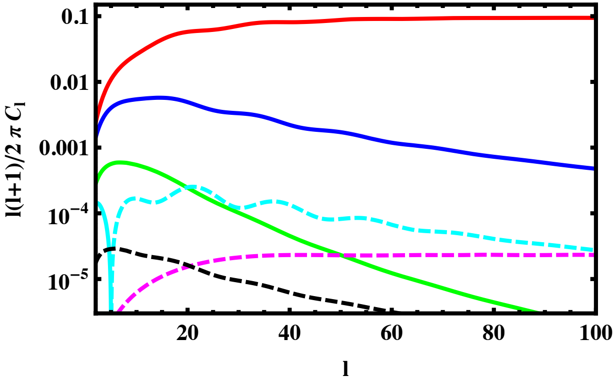
$$\cos\theta = \mathbf{n} \cdot \mathbf{n}'$$

The transverse power spectrum, $z' = z$ (from Bonvin & RD '11)



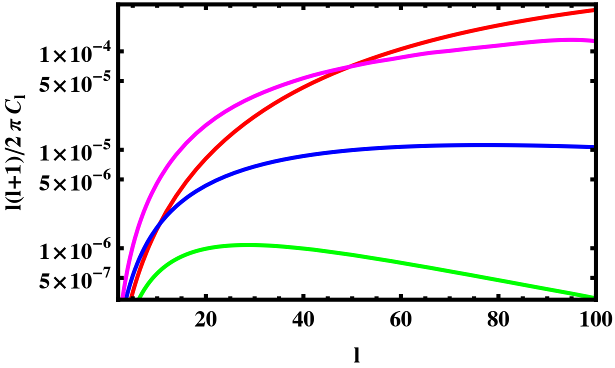
$z = 0.1$, $z = 0.5$, $z = 1$ and $z = 3$.

Contributions to the transverse power spectrum at redshift $z = 0.1, \ \Delta z = 0.01$ (from Bonvin & RD '11)



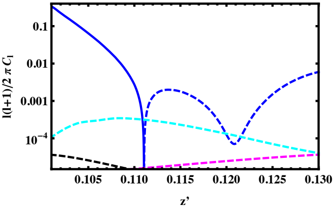
C_ℓ^{DD} (red), C_ℓ^{zz} (green), $2C_\ell^{Dz}$ (blue), $C_\ell^{Doppler}$ (cyan), $C_\ell^{lensing}$ (magenta) C_ℓ^{grav} (black).

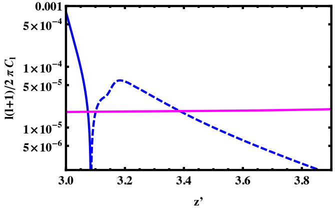
Contributions to the transverse power spectrum at redshift $z = 3$, $\Delta z = 0.3$ (from Bonvin & RD '11)

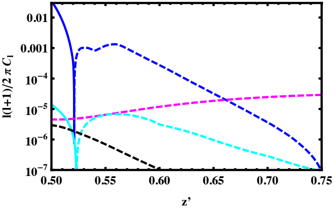


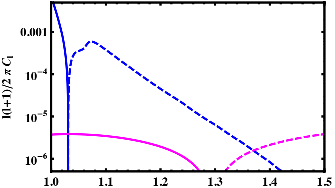
C_ℓ^{DD} (red), C_ℓ^{zz} (green), $2C_\ell^{Dz}$ (blue), $C_\ell^{lensing}$ (magenta).

The radial power spectrum





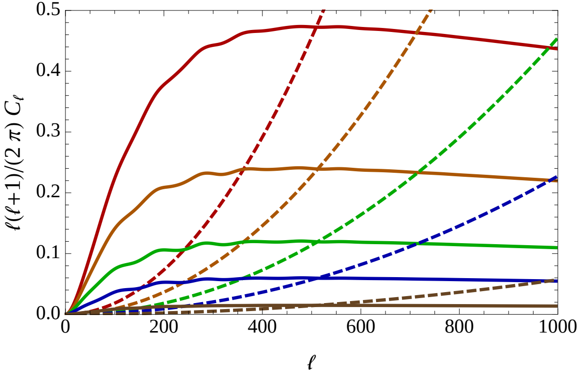




The radial power spectrum $C_\ell(z, z')$
for $\ell = 20$
Left, top to bottom: $z = 0.1,\ 0.5,\ 1,$
top right: $z = 3$

Standard terms (blue), $C_\ell^{lensing}$ (magenta), $C_\ell^{Doppler}$ (cyan), C_ℓ^{grav} (black),

Real experiments (DES): Shot noise vs. signal

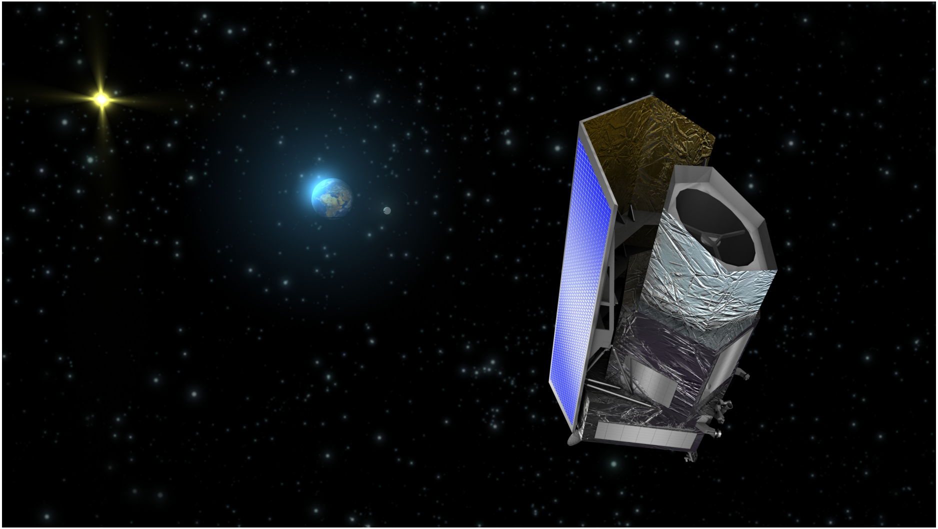


$\bar{z} = 0.55$
spectroscopic survey like DES
for shot-noise contribution.

(From Di Dio, Montanari, Lesgourgues, RD, 1307.1459 http://cosmology.unige.ch/tools)

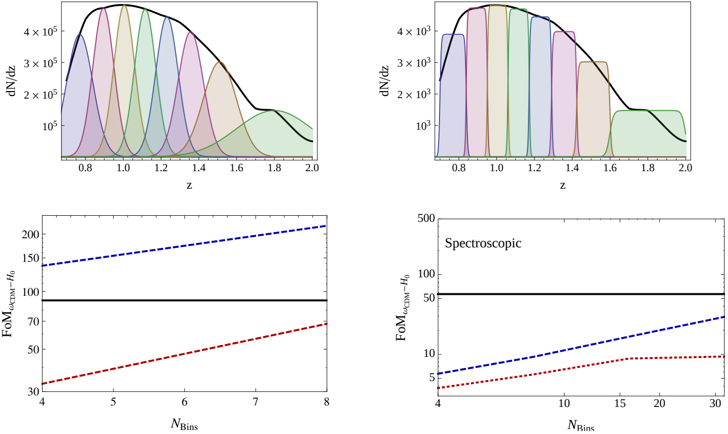
The angular power spectrum C_ℓ (solid lines) and the shot-noise contribution (dashed lines) for different top-hat window functions of half-widths: $\Delta z = 0.1$, $\Delta z = 0.025$, $\Delta z = 0.0125$, $\Delta z = 0.00625$, $\Delta z = 0.003125$.

$$C_\ell^{obs}(z,z) = C_\ell(z,z) + \frac{1}{N(z)}$$



(10^7 galaxy redshifts, 10^9 galaxies with photo-z)

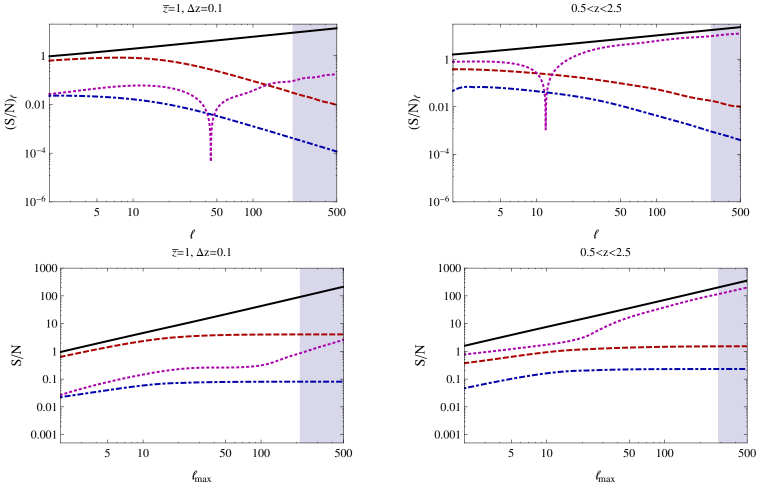
Real experiments (Euclid):



including cross-correlations, only auto-correlations

(From Di Dio, Montanari, RD, Lesgourgues, 1308.6186)

Real experiments (Euclid): Signal to Noise



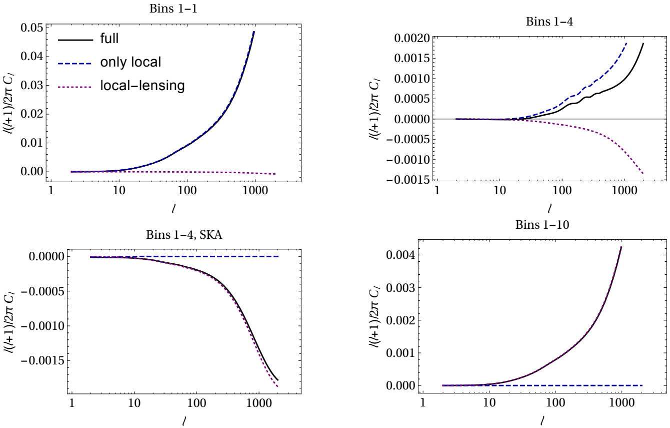
Signal to noise for different contributions:
density, redshift-space distortions, lensing, potential
(From Di Dio, Montanari, RD, Lesgourgues, 1308.6186)

Measuring the lensing potential

Well separated redshift bins measure mainly the lensing-density correlation:

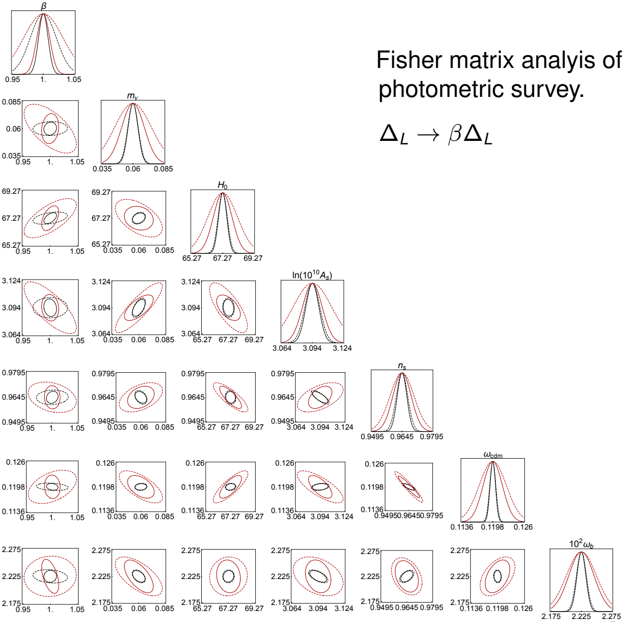
$$\langle \Delta(\mathbf{n}, z)\Delta(\mathbf{n}', z')\rangle \simeq \langle \Delta^L(\mathbf{n}, z)\delta(\mathbf{n}', z')\rangle \quad z > z'$$

$$\Delta^L(\mathbf{n}, z) = (2 - 5s(z)) \int_0^{r(z)} \frac{dr(r(z) - r)}{r(z)r} \Delta_2 \Psi(r\mathbf{n}, z)$$



(Montanari & RD) [1506.01369]

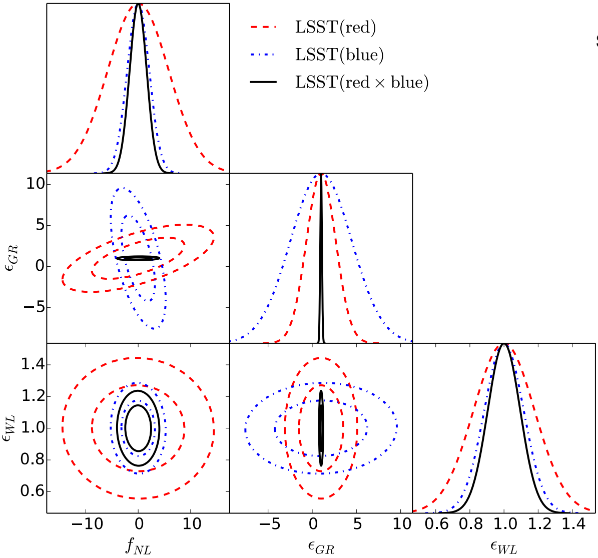
Testing GR with the lensing potential



Fisher matrix analyis of an Euclid-like photometric survey.

$$\Delta_L \to \beta\Delta_L$$

(Montanari & RD) [1506.01369]



standard parameters fixed

Alonso & Ferreira [1507.03550]

2nd order number counts

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D. Bertacca, R. Maartens, and C. Clarkson,[1405.4403,1406.0319]

J. Yoo and M. Zaldarriaga [1406.4140]

E. Di Dio, G. Marozzi, F. Montanari & RD [1407.0376]

2nd order number counts

The dominant terms are ($\propto (k/\mathcal{H})^4\Psi^2$)
(Di Dio, Marozzi, Montanari & RD, in preparation)

$$
\begin{aligned}
\Delta^{(2)Leading}(\mathbf{n},z) \quad \simeq \quad & \delta^{(2)} + \mathcal{H}^{-1}\partial_r^2 v^{(2)} - 2\kappa^{(2)} + \mathcal{H}^{-2}\left(\partial_r^2 v\right)^2 + \mathcal{H}^{-2}\partial_r v \partial_r^3 v \\
& + \mathcal{H}^{-1}\left(\partial_r v \partial_r \delta^{(1)} + \partial_r^2 v\, \delta^{(1)}\right) - 2\delta^{(1)}\kappa^{(1)} + \nabla_a \delta^{(1)} \nabla^a \psi \\
& + \mathcal{H}^{-1}\left(-2\partial_r^2 v\, \kappa^{(1)} + \nabla_a \partial_r^2 v \nabla^a \psi\right) + 2\left(\kappa^{(1)}\right)^2 - 2\nabla_b \kappa^{(1)} \nabla^b \psi \\
& - \frac{2}{r(z)}\int_0^{r(z)} dr \frac{r(z)-r}{r}\Delta_2\left(\nabla^b \Psi_1 \nabla_b \Psi_1\right) - 4\int_0^{r(z)} \frac{dr}{r}\nabla^a \Psi_1 \nabla_a \kappa^{(1)} .
\end{aligned}
$$

$$
\Delta^{(1)Leading} = \delta_\rho^{(1)} + \frac{1}{\mathcal{H}_s}\partial_r^2 v^{(1)} - 2\kappa^{(1)}
$$

$$
\psi^{(1)} = -2\int_0^{r(z)} dr \frac{r - r(z)}{r(z) r}\Psi\,, \qquad \kappa^{(1)} = -\Delta_2 \psi^{(1)}
$$

$$
\Psi^{(1)} = \frac{1}{r(z)}\int_0^{r(z)} dr \Psi
$$

The bispectrum

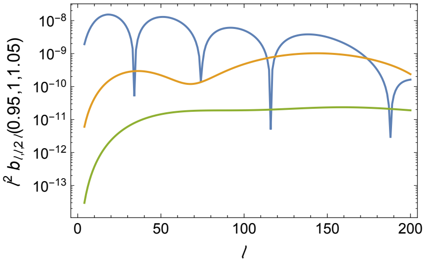
$$B\left(\mathbf{n}_1, \mathbf{n}_2, \mathbf{n}_3, z_1, z_2, z_3\right) = \left\langle \Delta\left(\mathbf{n}_1, z_1\right) \Delta\left(\mathbf{n}_2, z_2\right) \Delta\left(\mathbf{n}_3, z_3\right) \right\rangle$$

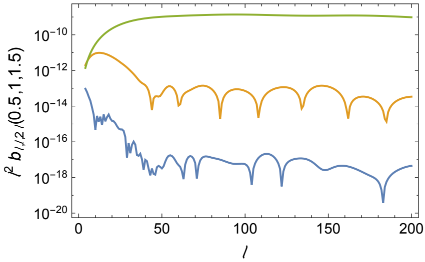
Statistical isotropy requires that

$$B^{m_1 m_2 m_3}_{\ell_1 \ell_2 \ell_3}(z_1, z_2, z_3) = \mathcal{G}^{m_1, m_2, m_3}_{\ell_1, \ell_2, \ell_3} b_{\ell_1, \ell_2, \ell_3}(z_1, z_2, z_3)\,,$$

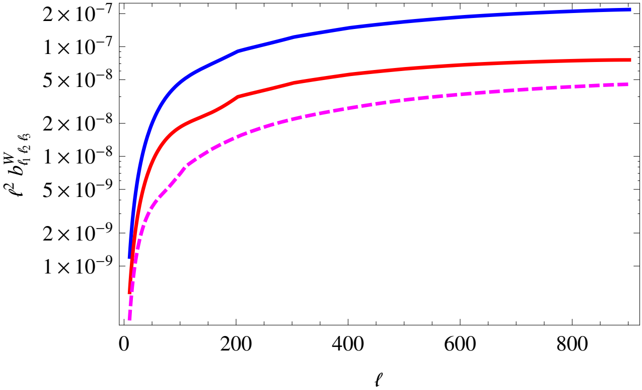
where $\mathcal{G}^{m_1, m_2, m_3}_{\ell_1, \ell_2, \ell_3}$ is the Gaunt integral.

The bispectrum





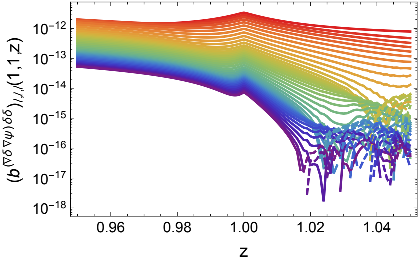
(Di Dio, RD, Marozzi & Montanari, in prep.)

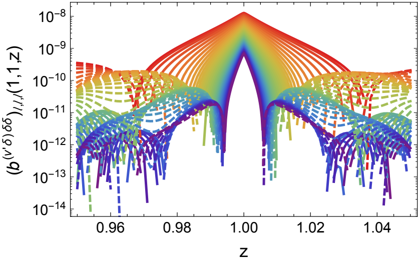


(from Di Dio, Marozzi, Montanari RD, '14) integrated from $z_{\min} = 0.2$ to $z_{\max} = 3$, for fixed $\ell_3 = 3$ while varying $\ell = \ell_1 = \ell_2$.

Density (blue), redshift space distortions (red), lensing (magenta)

The bispectrum





(Figures by Enea Di Dio)

Conclusions

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  $$r = \sqrt{r(z)^2 + r(z')^2 - 2r(z)r(z')\cos\theta}\ .$$
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- These spectra are not only sensitive to the matter distribution (density) but also to the velocity via (redshift space distortions) and to the perturbations of spacetime geometry (lensing) .

The spectra $C_\ell(z,z')$ and $b_{\ell_1,\ell_2,\ell_2}(z_1,z_2,z_3)$ depend sensitively and in several different ways on dark energy (growth factor, distance redshift relation), on the matter and baryon densities, bias, etc. Their measurements provide a new route to estimate cosmological parameters and to test general relativity.