AugerPrime:
the upgrade of the Pierre Auger Observatory

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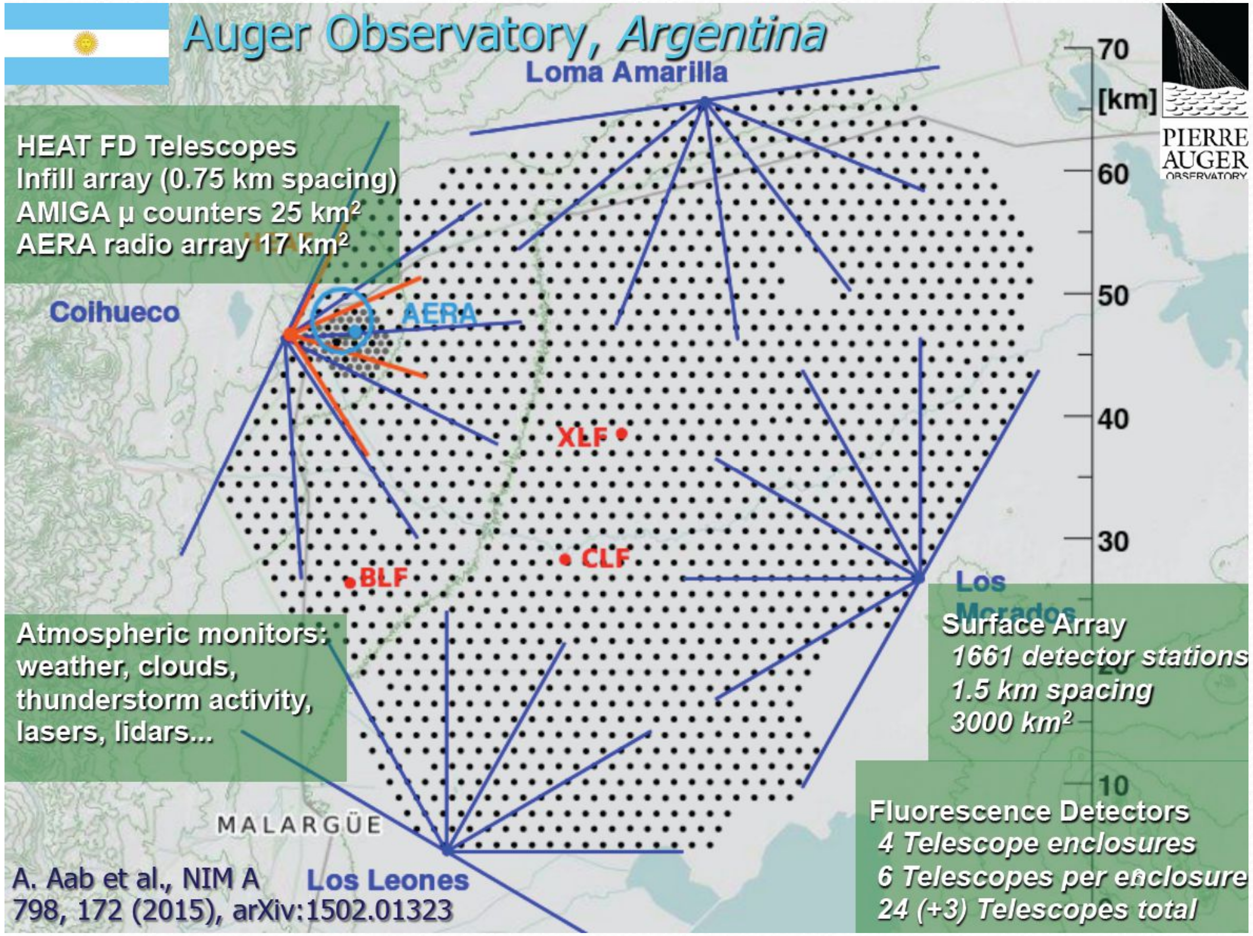
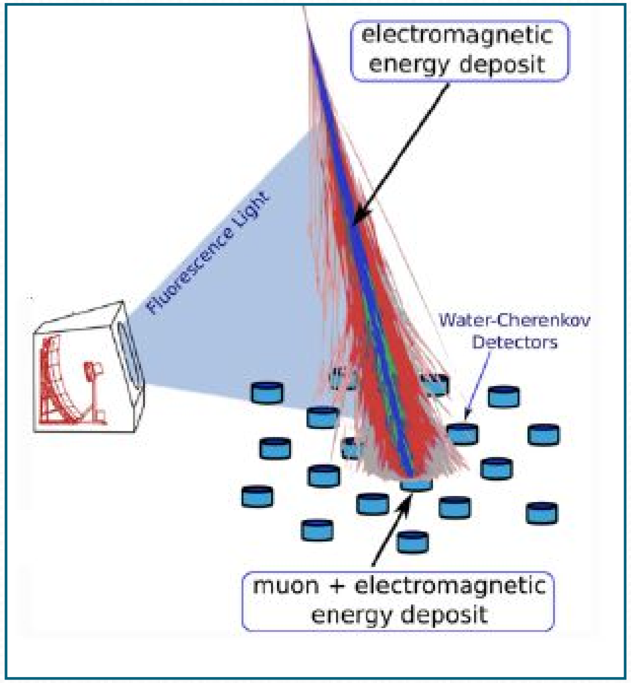
*INFN Lecce

Pierre Auger Observatory: Hybrid Detector

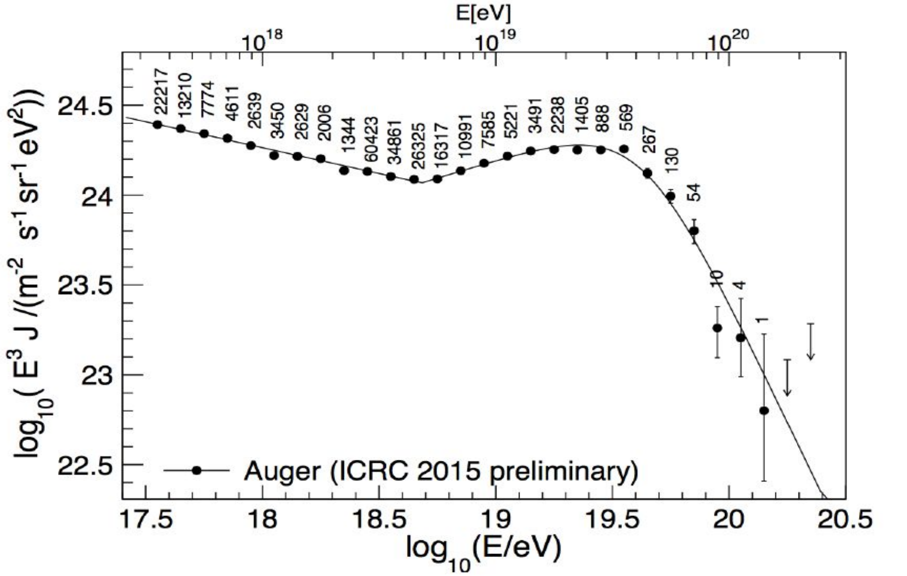
(refer to Isabelle Lhenry-Yvon talk on Thu 06/07)



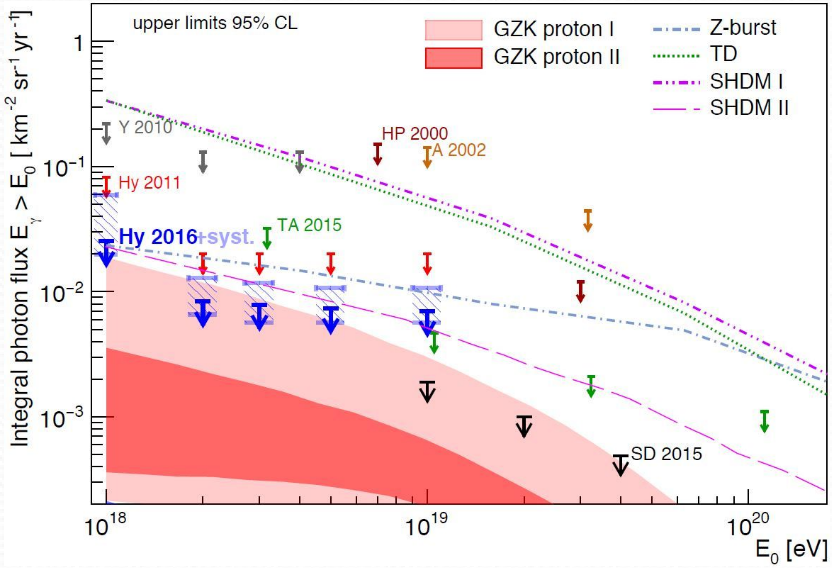


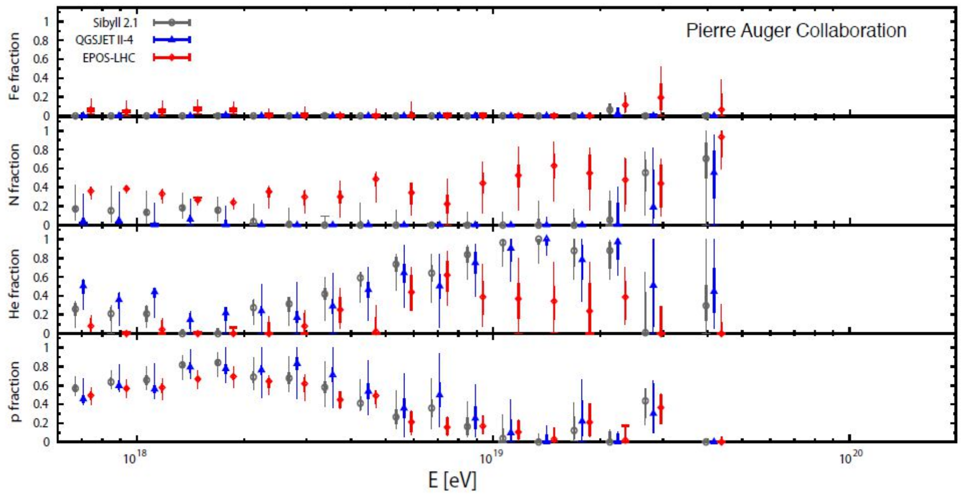
Scientific results from the PAO

Strong flux suppression above 5x10[19]eV



Photon and neutrino upper limits
(exclusion of top-down models)





Mass composition change
(but no data above 5 x 10[19] eV)

The upgrade Science Case

The data collected after 2017 must provide additional measurements to allow us to address the following key objective:

1. The mass composition and the origin of flux suppression at the highest energies
   - Understanding the origin of the flux suppression will provide fundamental constraints on the astrophysical sources and will allow a more reliable estimates of neutrino and gamma-ray fluxes at UHE.
2. Proton contribution in the flux suppression region ($E > 5\text{x}\,10^{19}$ eV)
   - Estimate the physics potential of existing and future cosmic ray, neutrino and gamma-ray detectors; Search of proton astronomy
3. Fundamental particle physics at energies beyond reach of man-made accelerators.
   - Study extensive air showers and hadronic multiparticle production.

Mass composition measurement above 5×10^{19} eV with a sensitivity to the proton flux as small as 10%.

The strategy

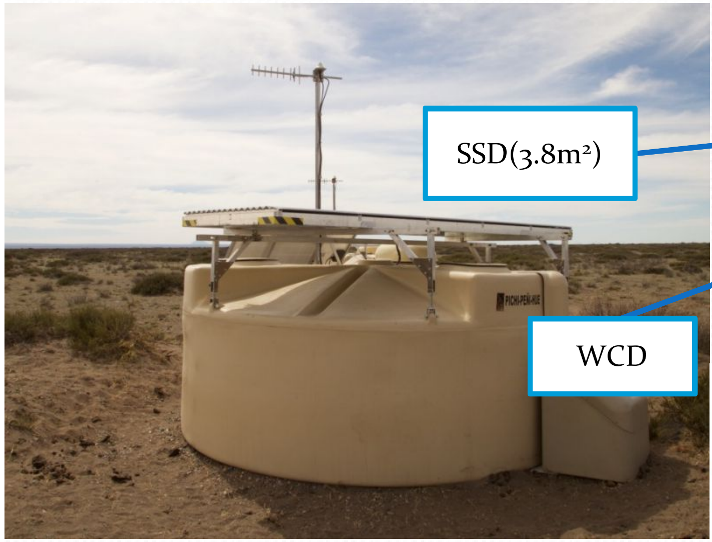
Measure with the Pierre Auger Observatory until 2025.

(MOUs have been signed in Nov 2015)

Main improvements:

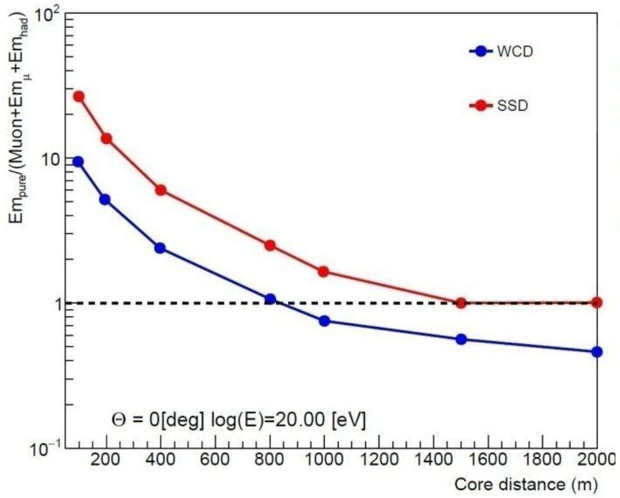
1. A new detector above each of the existing water-Cherenkov detectors (WCD)
2. A new electronics for the SD and the extension of the dynamic range
3. Extended FD operation
4. Underground Muon Detector with AMIGA to have a direct muon measurement

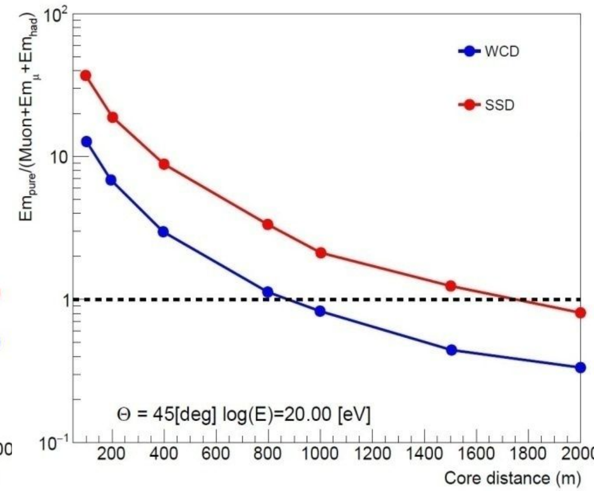
Scintillator Surface Detector Measurement

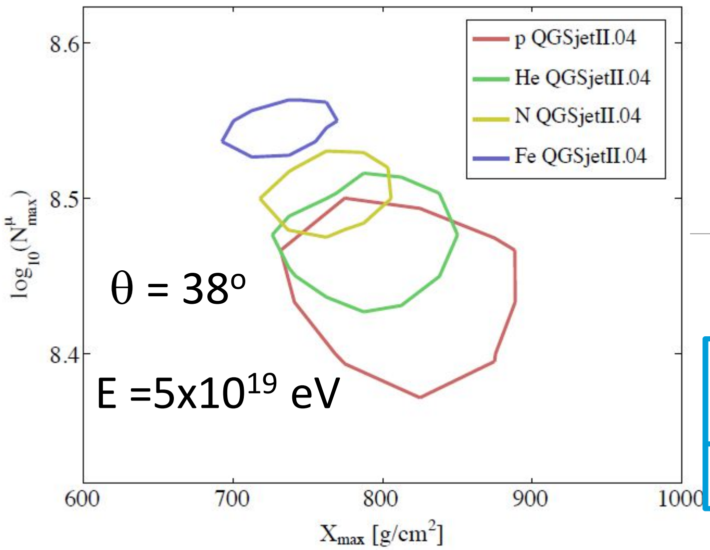


100% duty cycle

Complementarity of particle response used to discriminate electromagnetic and muonic components of air showers





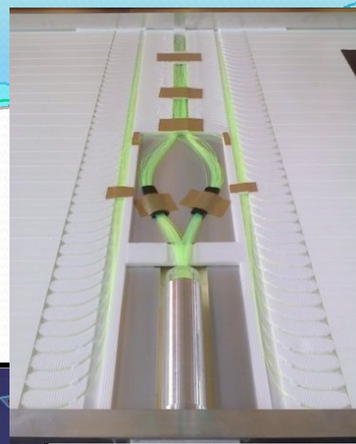


1σ contour of the number of muons at maximum of the muon shower development

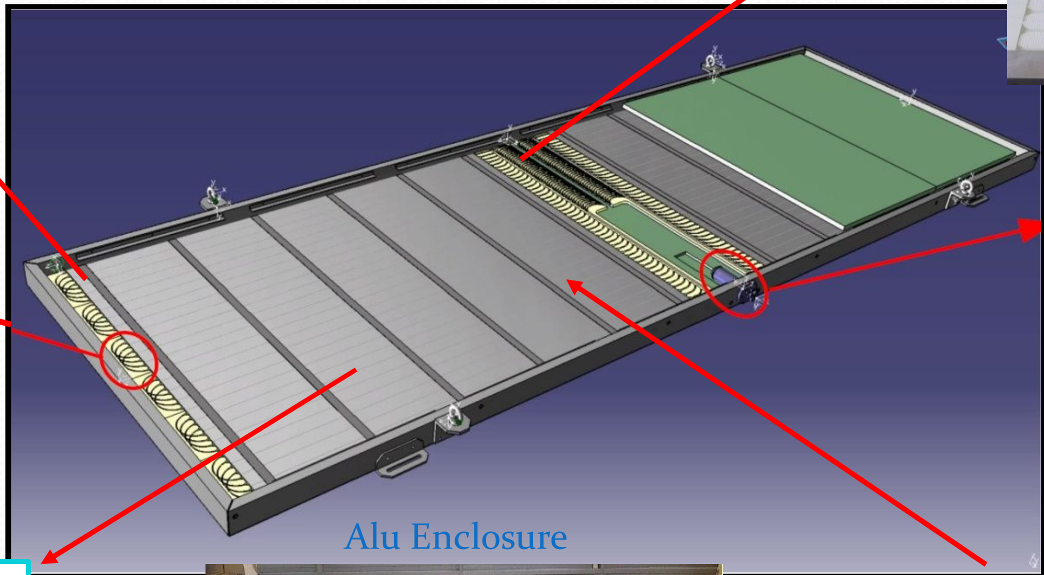
Measure muon component for composition

SSD:The detector

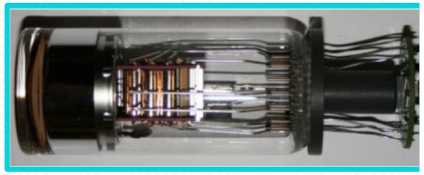
WLS
fibers+routers



Extruded Scintillator bars with 2 holes

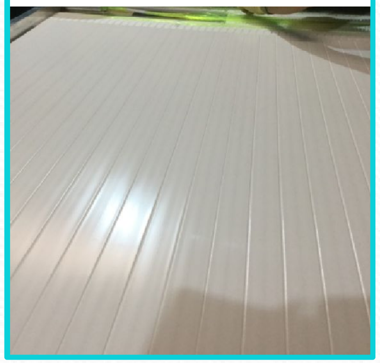


WLS
fibers+routers

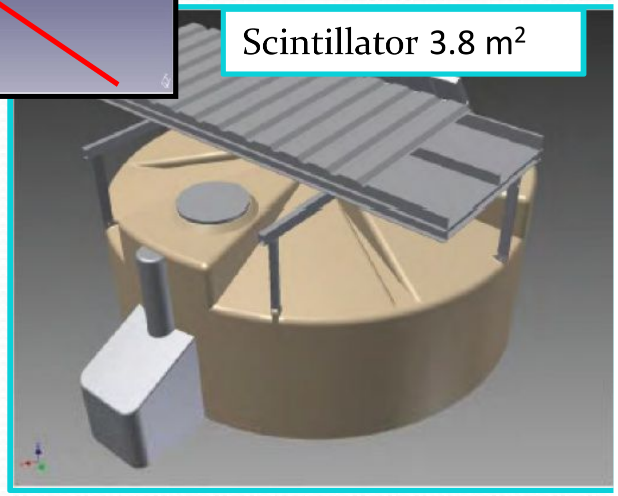


PMT

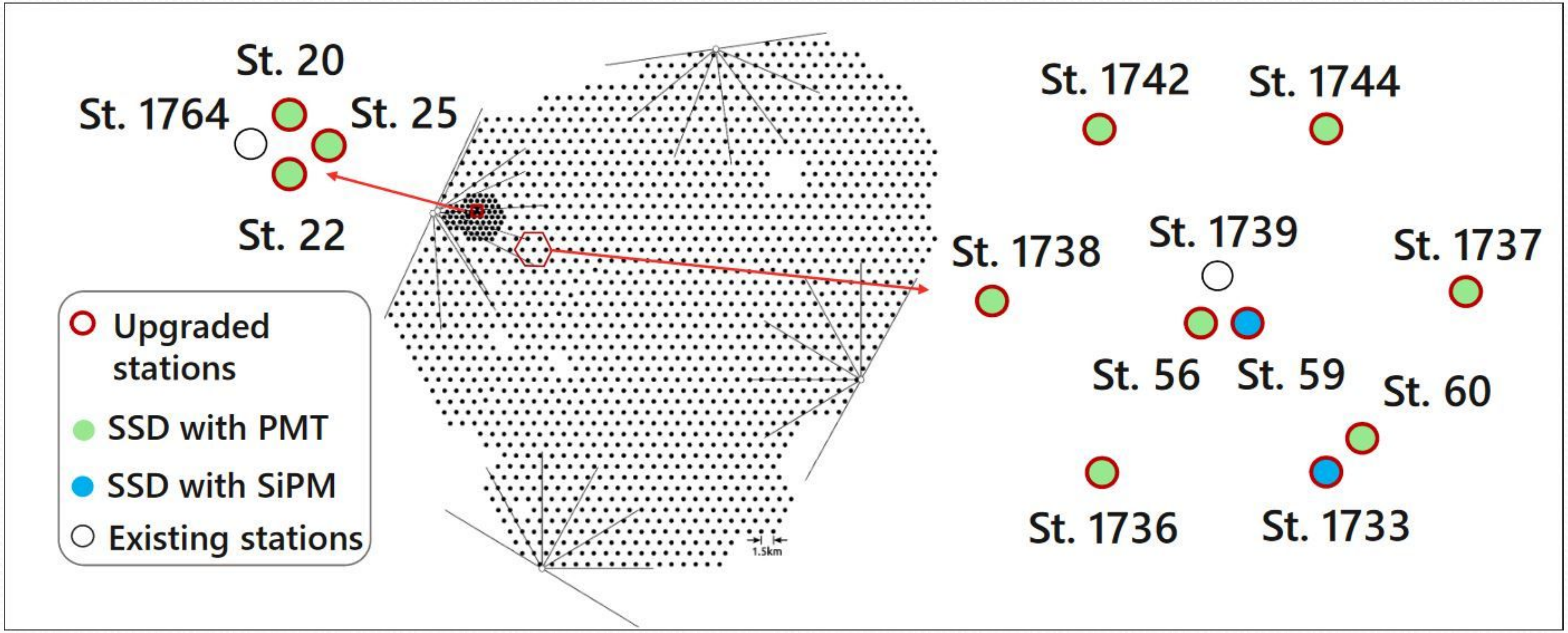
Extruded scintillator bars
160cm long



Scintillator 3.8 m^2



SSD:The Engineering Array

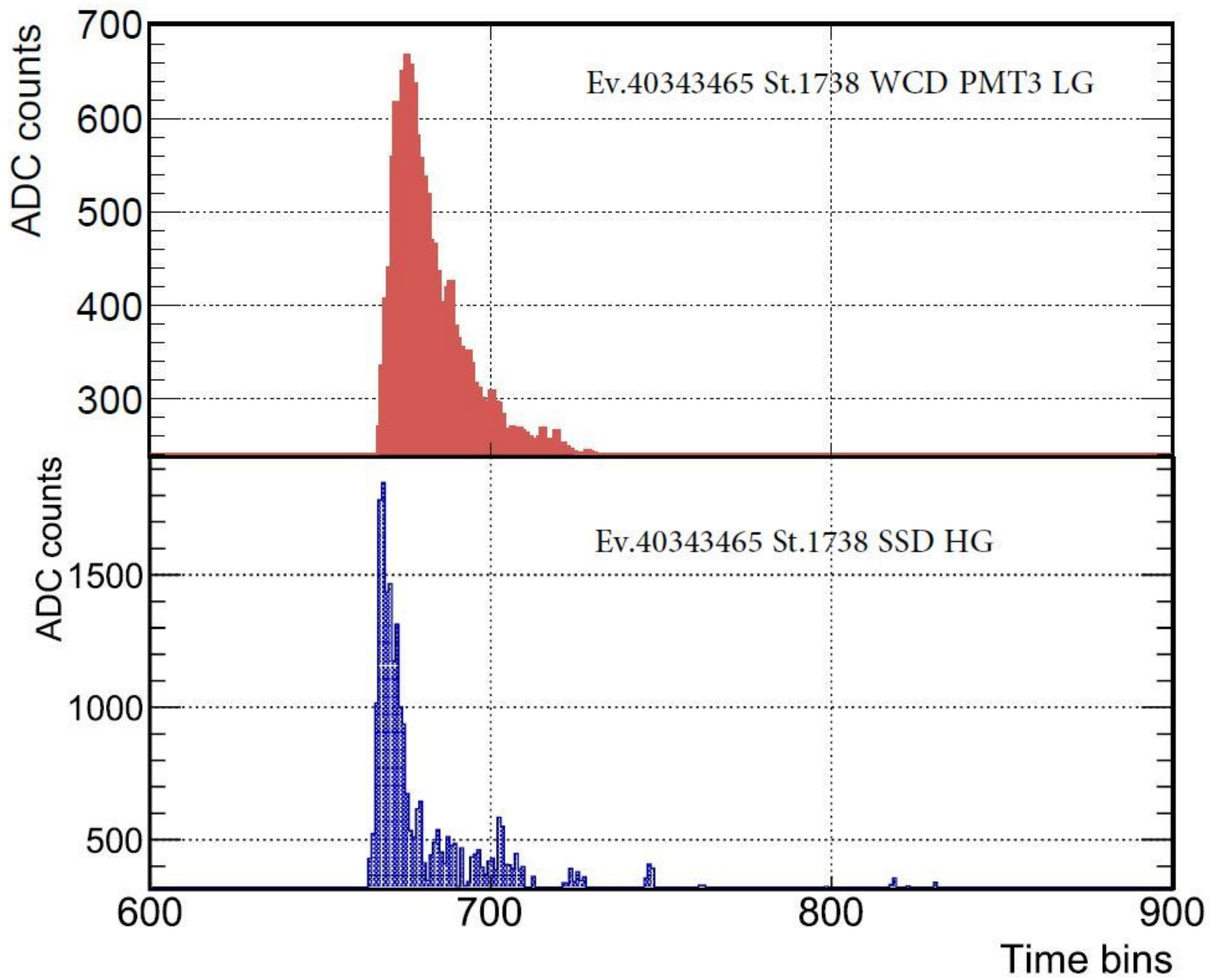


An engineering array consisting of 12 AugerPrime detector stations has been in operation since Oct. 2016 in the Observatory near Malargüe, Argentina

SSD:The Engineering Array

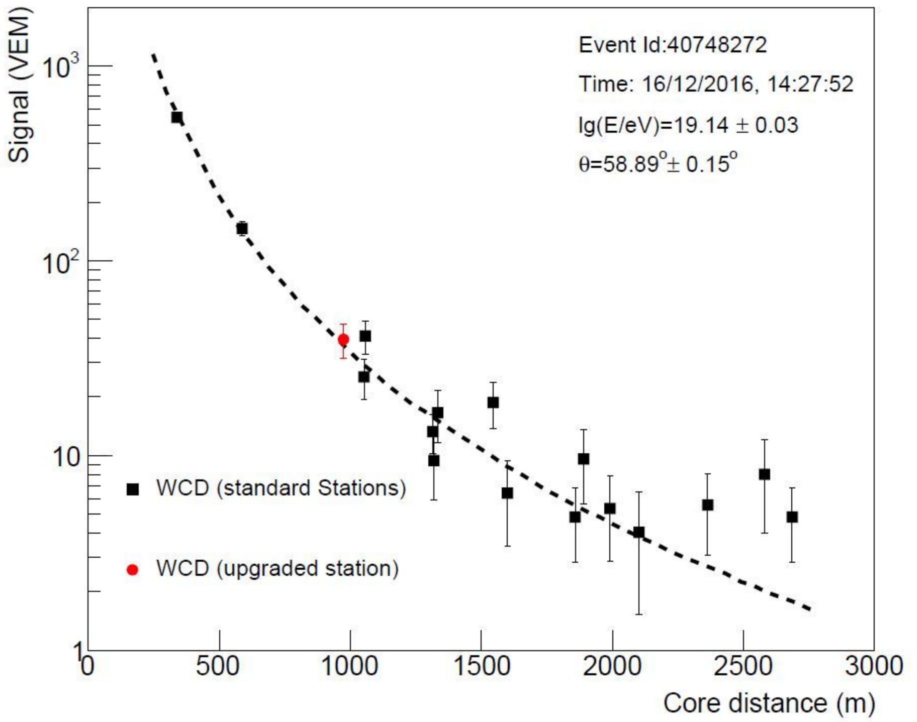
The detection of shower events in each SSD station is triggered by the WCD

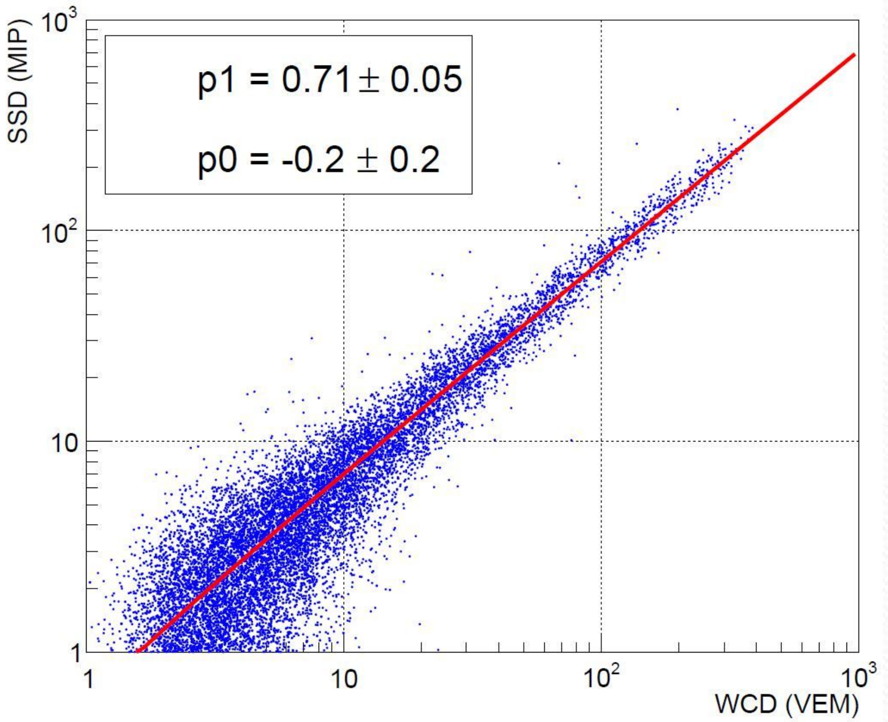
The two detector types have different responses to the electromagnetic (EM) and muonic components of the EAS.



The WCD and SSD signals for the same event

SSD:The Engineering Array

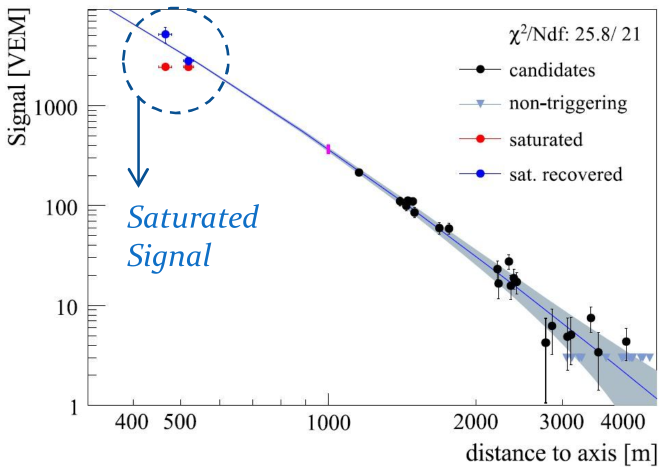




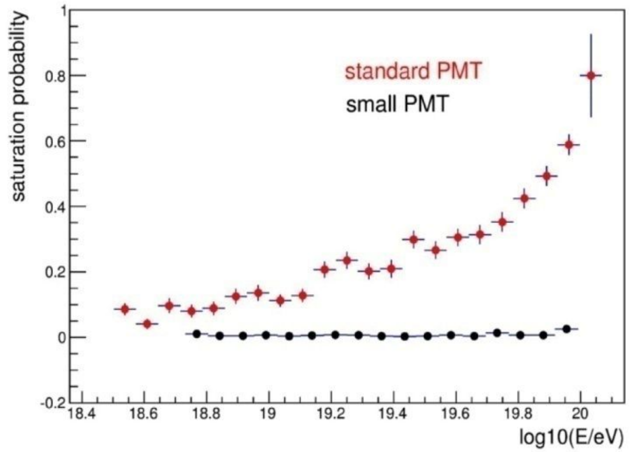
One event reconstructed with the regular 1500 m array in close proximity to the EA. The reconstructed signals in the EA are compared with the LDF of the event

Correlation of the signals of the SSD and the WCD. Both signals have been calibrated

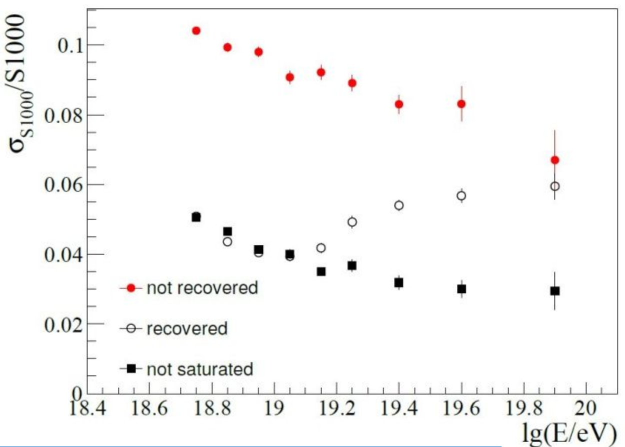
SD Electronics: small PMT



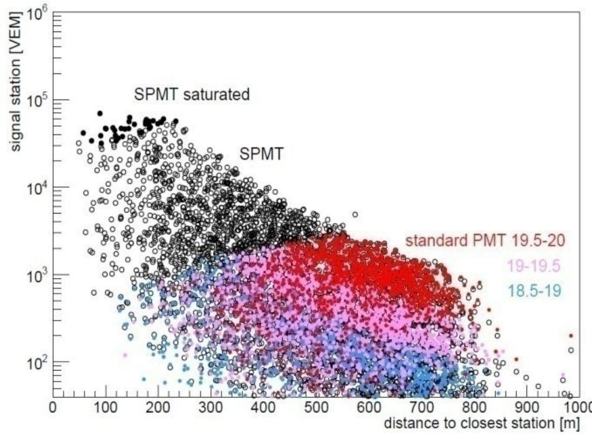
Lateral Distribution of the signal sizes recorded in SD detectors



Probability of having at least one saturated station in an event as function of energy, obtained from simulation for standard and small PMT

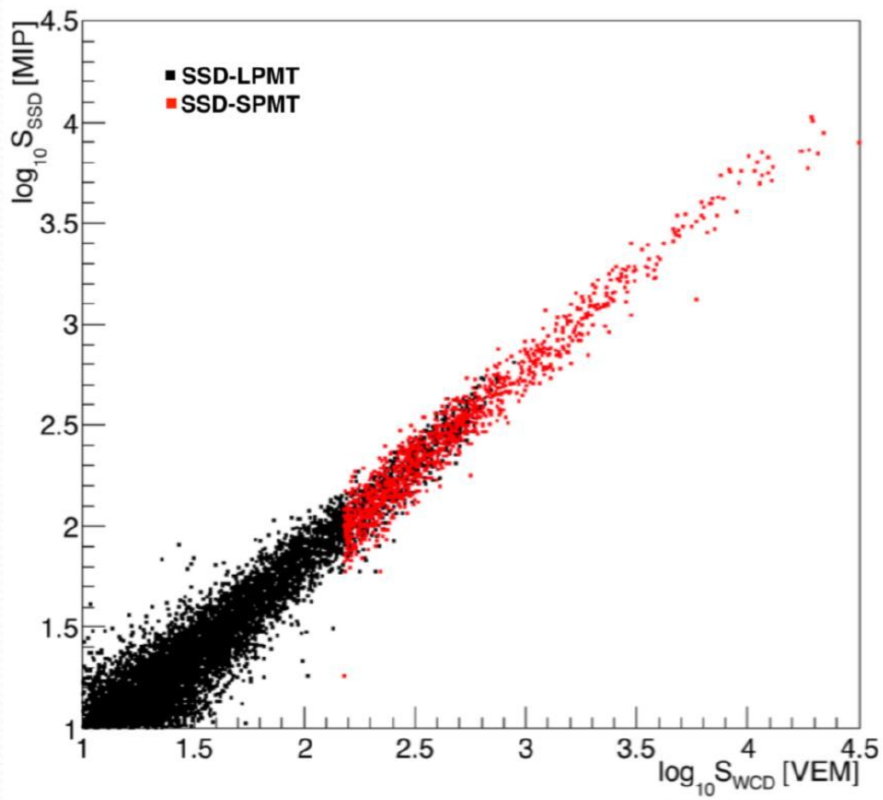


Resolution of the Reconstructed S(1000)

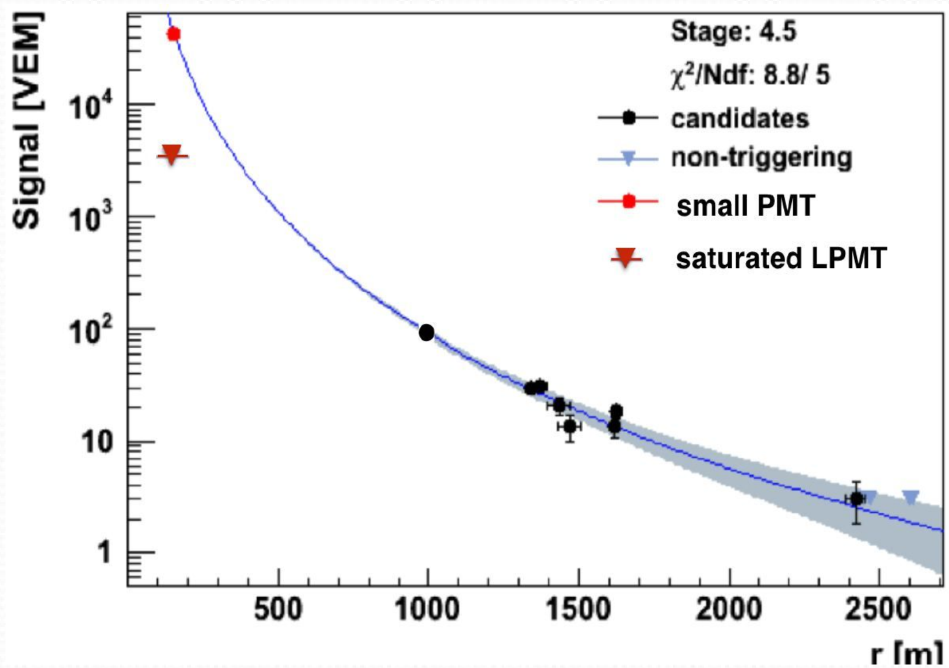


The distribution of the expected signals as a function of the distance between the shower axis and the closest SD station

SD Electronics: small PMT

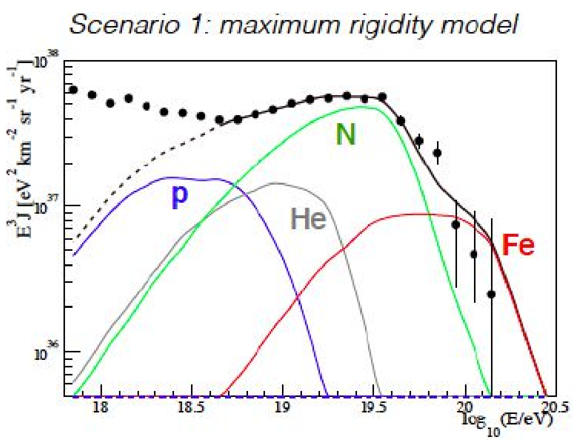


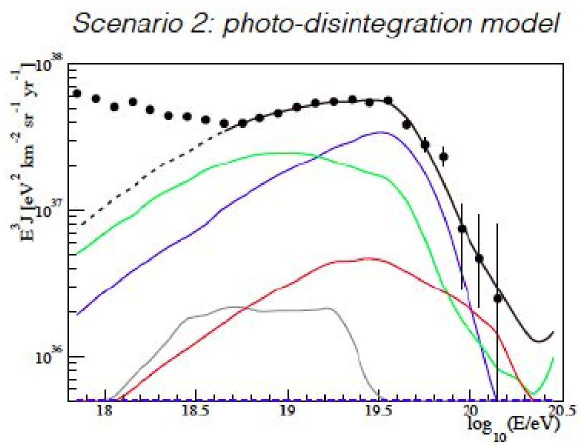
Relation between the signals in one of the water-Cherenkov stations and in the corresponding scintillator as measured by one detector of AugerPrime.

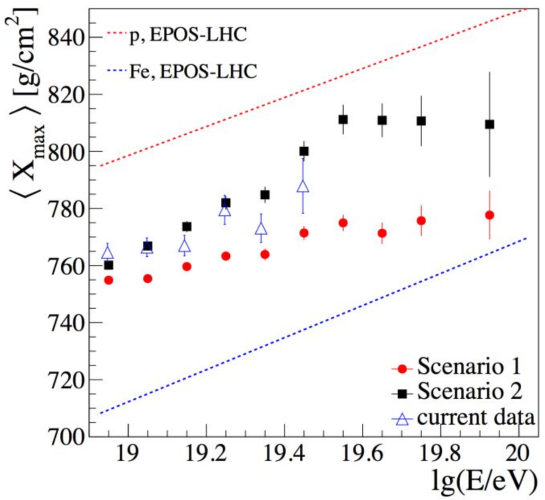


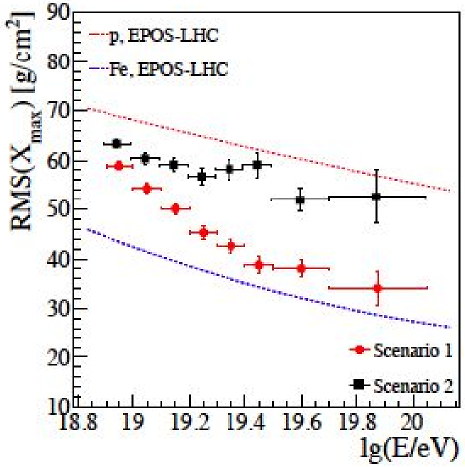
Lateral distribution for one event measured in the engineering array of AugerPrime. The signal in the station closest to the shower core (153 m) is recorded by the SPMT the signal in the LPMTs is saturated (red triangle)

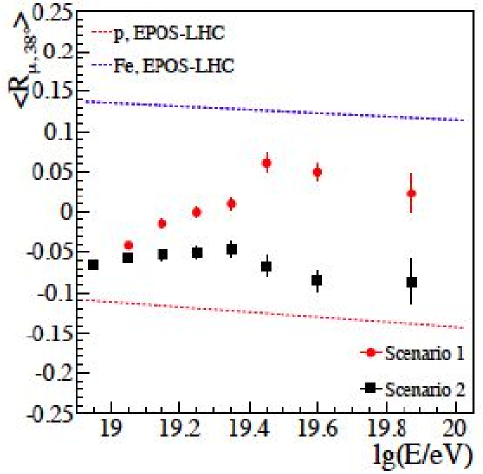
Science Impact of upgrade





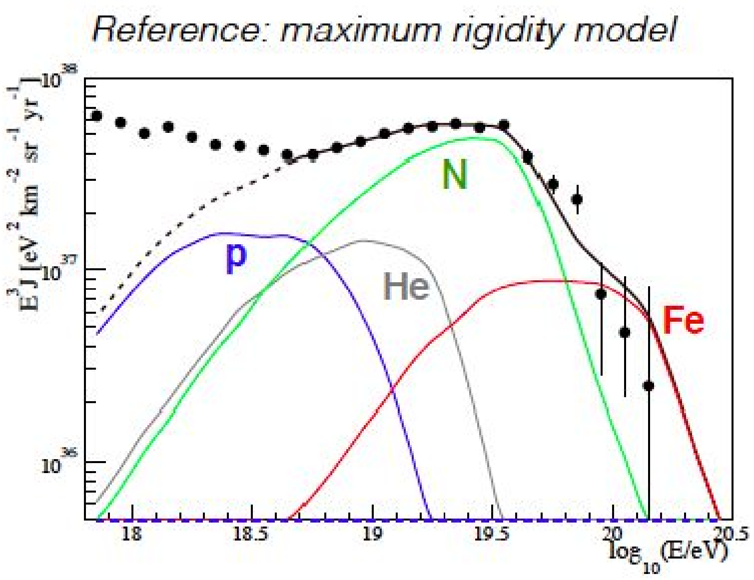


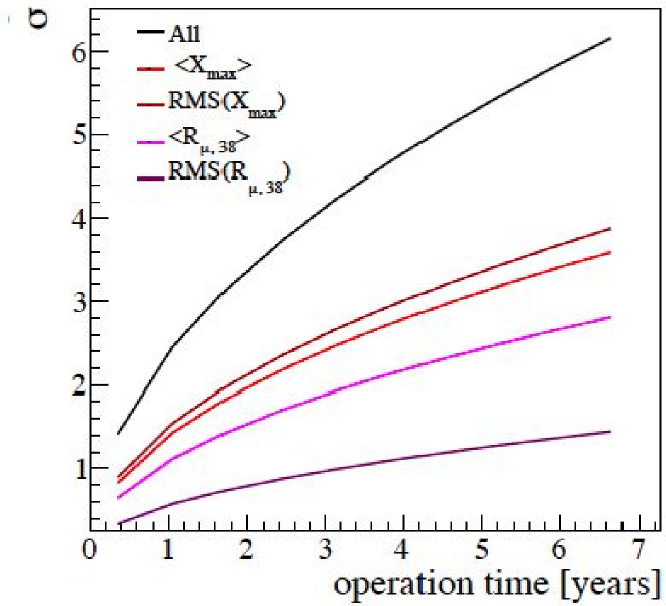




Science Impact of upgrade

- Physics reach: detection of 10 % proton contribution
- Significance of distinguishing scenarios with and without 10% of protons





•Standard scenario 1 (almost no protons)
• Scenario 1 with 10 % protons added

Timeline for new SDE and SSD

- July 2016: Engineering Array (12 stations) ready!
- October 2016: Engineering Array taking data!
- Nov 2016-July2017: Evaluation of detectors
- September 2017 Detector construction start.
- Start the deployment in 2018
- Till 2025: Data taking (up to 40,000 km^2 sr yr)
- Similar event statistics as collected so far will be reached with upgraded detectors.

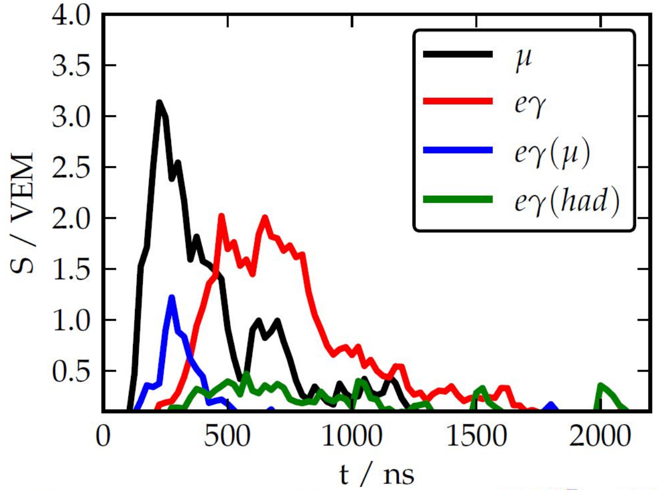
Summary and Outlook

- The AugerPrime engineering array has been taking data since October 2016
  - EA stations have good stability in operation.
  - WCDs and SSDs of EA stations perform well in shower detection.
- The dynamic range of the AugerPrime upgrade of the Pierre Auger Observatory has been extended to particle densities as high as few thousand per m^2, thus allowing us to measure full signals from all the stations of the air shower footprints at the ground down to a distance of about 250 m from the shower core.
- AugerPrime will allow a study of mass composition above 5×10^{19} eV and address:
  - Origin of the flux suppression (GZK energy loss Vs. maximum energy of sources)
  - Proton contribution of more than 10% above 5×10^{19} eV?

    (particle astronomy, GZK γ and ν fluxes → future experiments)
  - New particle physics beyond the reach of LHC?

Backup slides

SSD Measurement: Universality approach



The shower universality method **predicts for the** entire range of **primary masses the air-shower** characteristics on the ground **using only three** parameters: ***E, Xmax and Nμ***

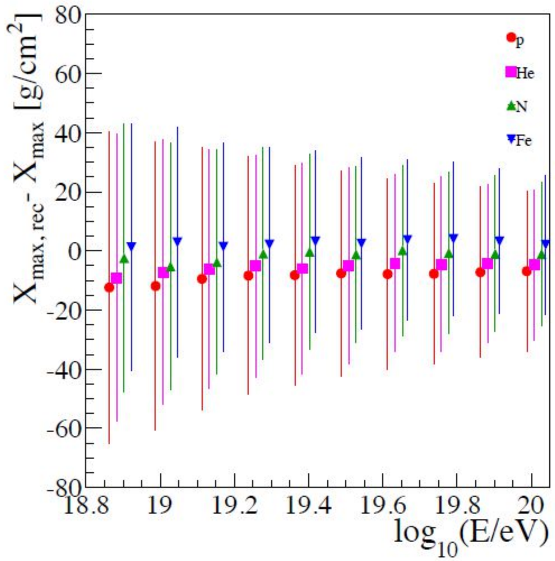
The parameter could be **estimated from the integrated signal and the temporal structure** of the signal measured in individual stations. Event by event basis

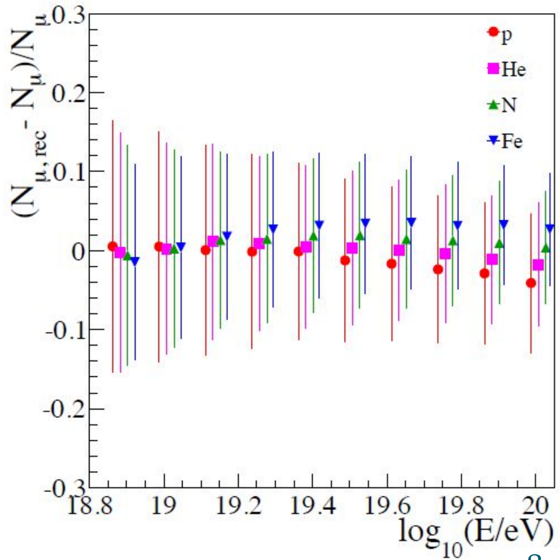
$$S_{\text{tot}} = S_{\text{em}}(r, DX, E) + N_{\mu}^{\text{rel}} \left[S_{\mu}^{\text{ref}}(r, DX, E) + S_{\text{em}}^{\mu}(r, DX, E) \right] + (N_{\mu}^{\text{rel}})^{\alpha} S_{\text{em}}^{\text{low-energy}}(r, DX, E)$$

Applying the Universality method it is possible to take into account the **correlation between the WCD and the SSD. The parameters now are more (**Xμmax ,Xmax, Nμ) in the model.

This allows a measurement of the **number of muons on a event by event basis and the relation** between Xμmax ,Xmax and Nμrel can be calibrated.

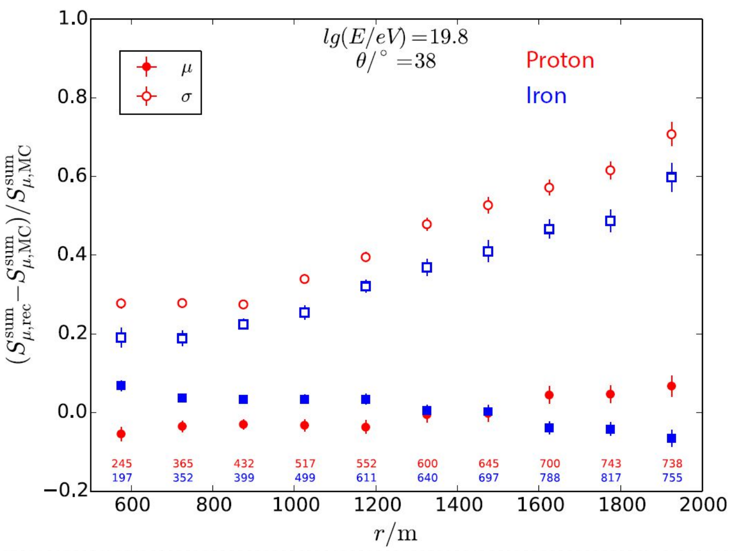
The **resolutions of the method are obtained from parameterizations and interpolations of EPOS-LHC** simulations at fixed energies and zenith angles and are shown for events up to 60º.





SSD: Matrix Inversion Method

Single Station Analysis

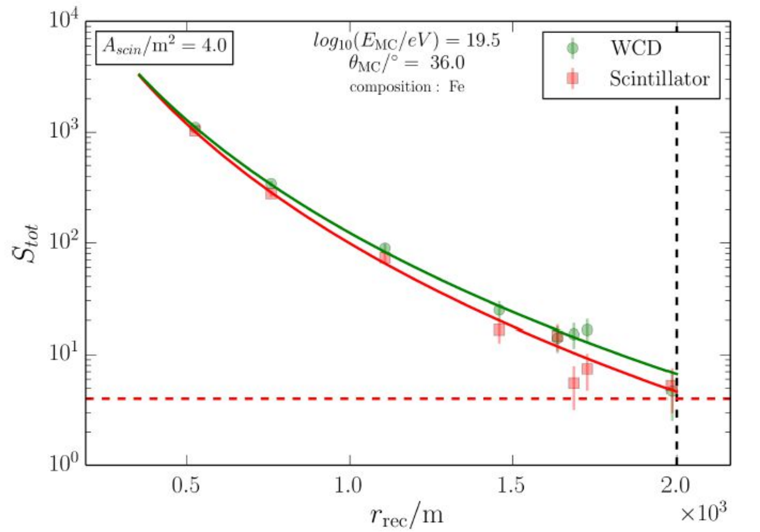


$$S_{\mu,\text{WCD}} = a\,S_{\text{WCD}} + b\,S_{\text{SSD}}$$

Lateral Distribution Analysis

A **parameterization of the LDF for the SSD** was done using simulation.

Simulated Fe LDFs fit for WCD and SSD



The **matrix inversion algorithm is then applied** to the LDF values for the WCD and SSD to calculate the **muonic signal expected in a WCD at 800 m core distance,** *Sμ(800).*

$$f_{p,Fe} = \frac{|\langle S_{Fe}\rangle - \langle S_p\rangle|}{\sqrt{\sigma(S_{Fe})^2 + \sigma(S_p)^2}} \sim 1.5$$

Figure of merit

SD Electronics

Auger electronics based on a 15 years old design

1. Increase of the data quality (better timing, dynamic range and μ identification):
   a) faster sampling of ADC traces (40 → 120 MHz)
   b) more precise absolute timing accuracy (new GPS receiver)
   c) increase the dynamic range by adding a 1" PMT (SD PMTs are 9") small PMT
2. Faster data processing and more sophisticated local triggers
   a) more powerful processor and FPGA
3. Improved calibration and monitoring capabilities
4. New components:
   1. Connection to the SSD and any additional (R&D) detectors
   2. Prolong lifetime and reduce failure rate

Can be swapped in place with old design
(same power communications, hardware interfaces...)

The Upgrade Unified Board (prototype)

Fluorescence Detector Operation

- The FD provides exceptional information about extensive air showers (model-independent energy reconstruction and direct measurement of the longitudinal development profiles)
- The main limitation of the FD is the duty cycle, currently at the level of 15%.

Increase the exposure for cosmic ray events above 10^{19} eV by extending the FD measurement into hours with high night sky background (NSB)

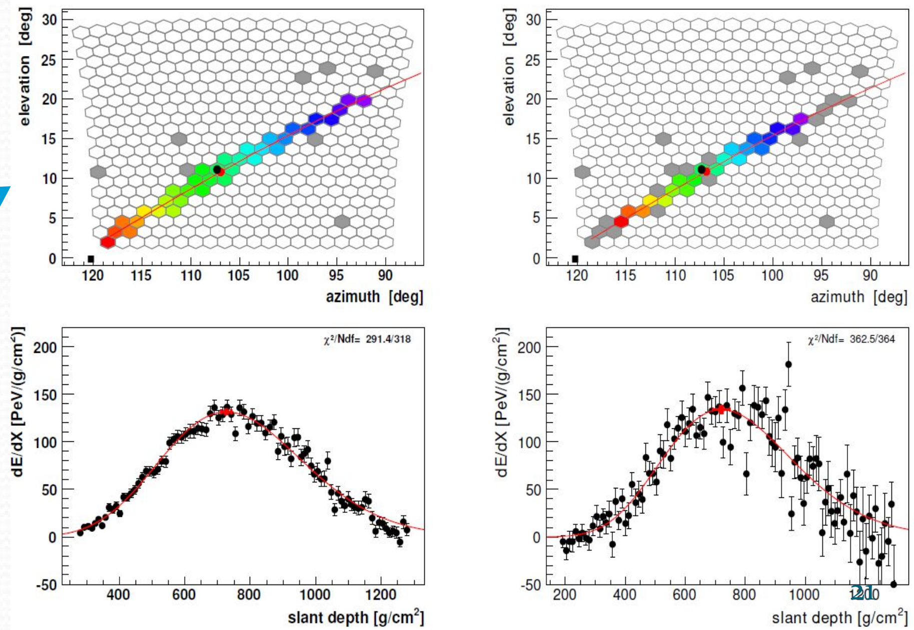
A significant increase of the duty cycle is possible by the extension of the FD operation to times at which a large fraction of the moon in the sky is illuminated. During such operations the PMT gain must be reduced (lower HV) to avoid an excessively high anode current.

10x reduced PMT gain by reducing supplied HV

satisfy linearity, stability and lifetime

Existing measured air showers have been analyzed with the standard reconstruction chain after adding random noise to the ADC traces.

keep very high selection efficiency and reconstruction.

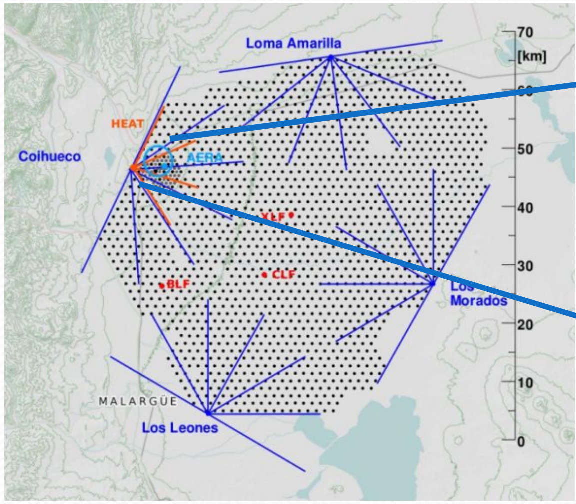


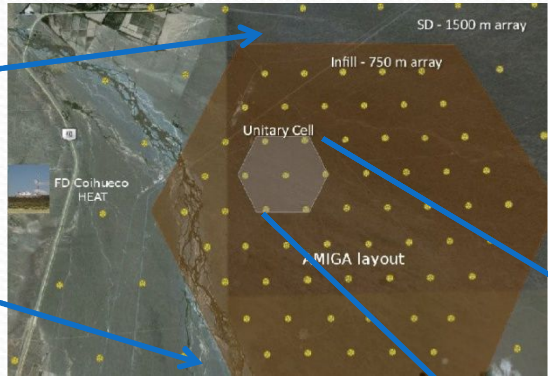
The Underground Muon Detector

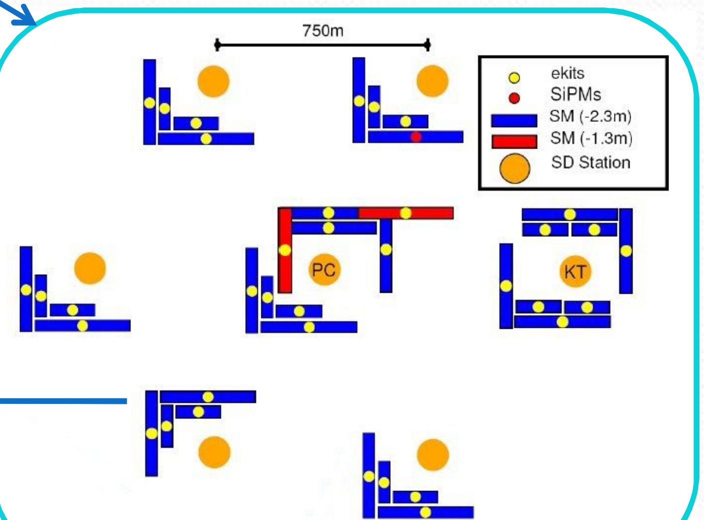
The UMD is required in the existing SD infill area of 23.5 km^2 in order to provide important direct measurements of the shower muon content and its time structure while serve as verification and fine-tuning of the methods used to extract muon information with the SSD and WCD measurements.

61 AMIGA muon detectors (30 m^2) are planned to be deployed on a 750m grid (a total area of 23.5 km^2)

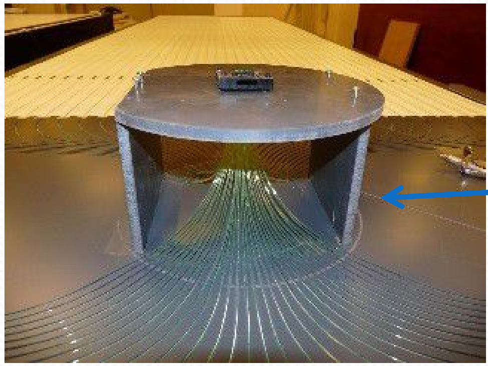
AMIGA layout: an infill of surface stations with an inter-detector spacing of 750 m. Plastic scintillators of 30m^2 are buried under 280 g/cm^2 of vertical mass to measure the muon component of the showers. The Unitary Cell indicates the prototype area of the muon detector.







AMIGA-Unitary cell

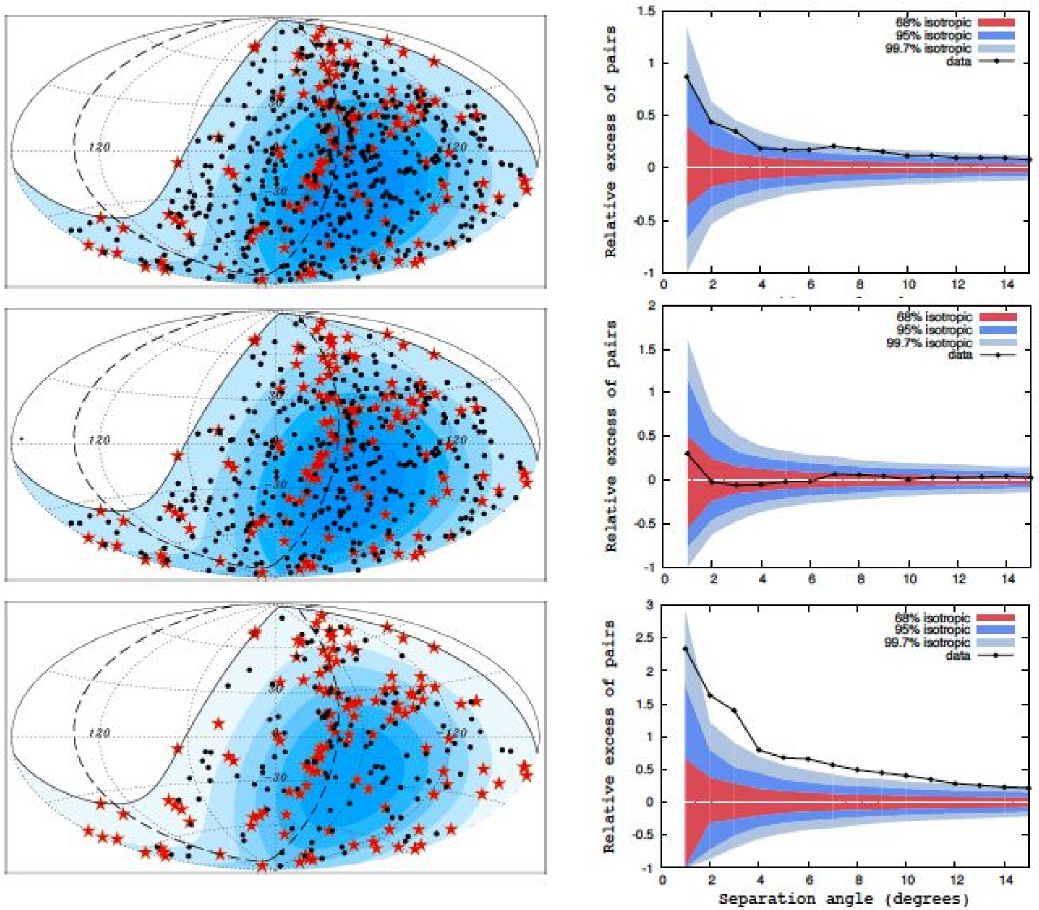


Science Impact of upgrade: composition-enhanced anisotropy

Modified Auger data set
(E > 4×10^{19} eV, 454 events)

X_{max} assignment according to maximum rigidity scenario

10% protons added, half of which from within 3° of AGNs



all 454 events

proton depleted data set (326)

proton enhanced data set (128)

Science impact of upgrade: photon and neutrino flux limits

Expected sensitivity on the flux of photons and neutrinos.
In addition to the conservative estimates based on the increase of statistics, also the projected photon sensitivity for the ideal case of being able to reject any hadronic background due to the upgraded surface detector array is shown.