Cosmic ray antiprotons : where are we ?

Pierre Salati – LAPTh & Université Savoie Mont Blanc

Outline

1) Cosmic rays as an indirect probe for dark matter

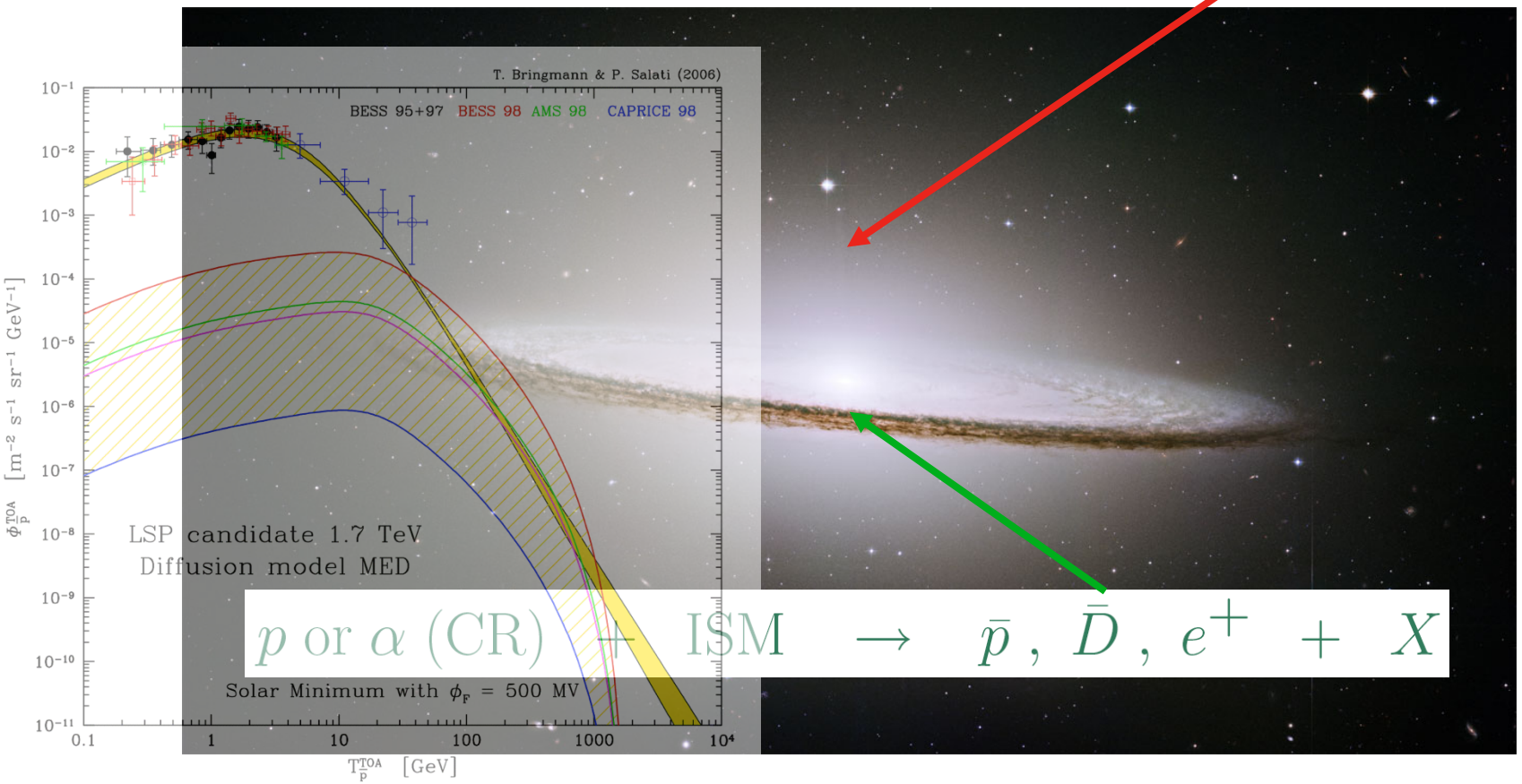
2) A new estimate of the antiproton background

3) Extracting limits on DM particle properties

1) Cosmic rays as an indirect probe for dark matter

• Indirect Detection – WIMPs continuously annihilate and produce SM particles such as gamma-rays, neutrinos, but also rare antimatter species like positrons, antiprotons and even antideuterons.

$$\chi + \chi \rightarrow q\bar{q},\, W^+W^-,\, \ldots \rightarrow \boxed{\gamma,\, \bar{p},\, \bar{D},\, e^+,\, \nu's}$$



AMS Collaboration

CERN, Geneva, 15 April 2015

"AMS Days at CERN" and Latest Results from the AMS Experiment on the International Space Station

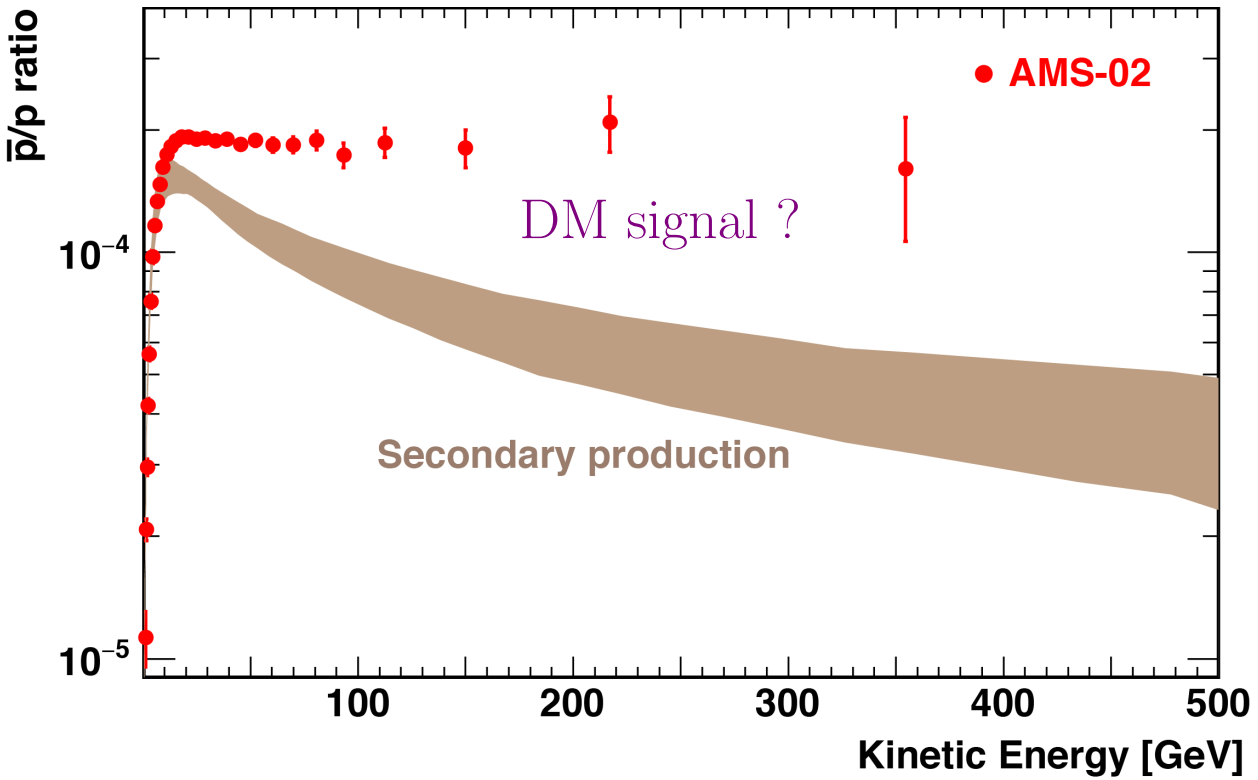


Figure 1. Antiproton to proton ratio measured by AMS. As seen, the measured ratio cannot be explained by existing models of secondary production.

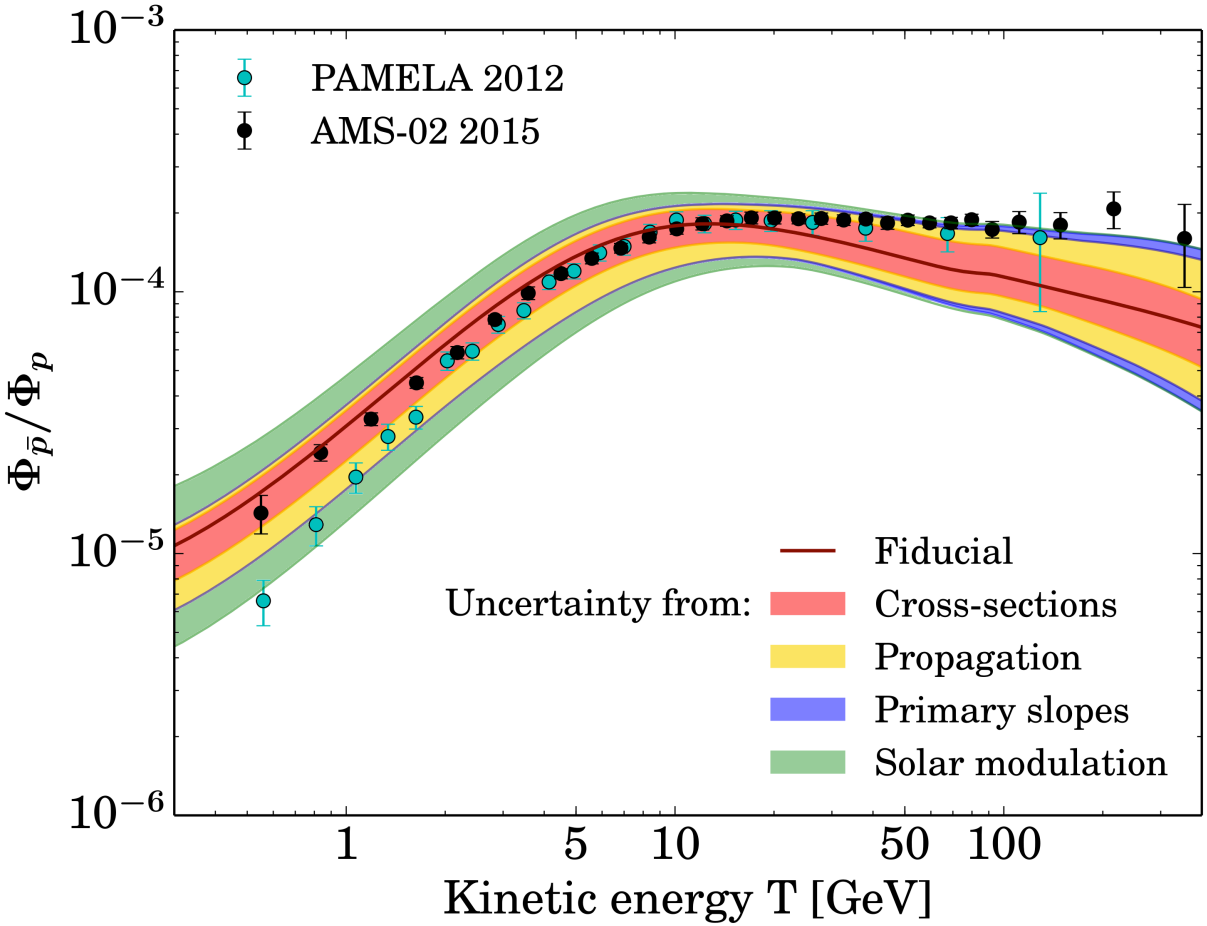
Has an antiproton excess been discovered ?

AMS Collaboration
CERN, Geneva, 15 April 2015

"AMS Days at CERN" and Latest Results from the AMS Experiment on the International Space Station

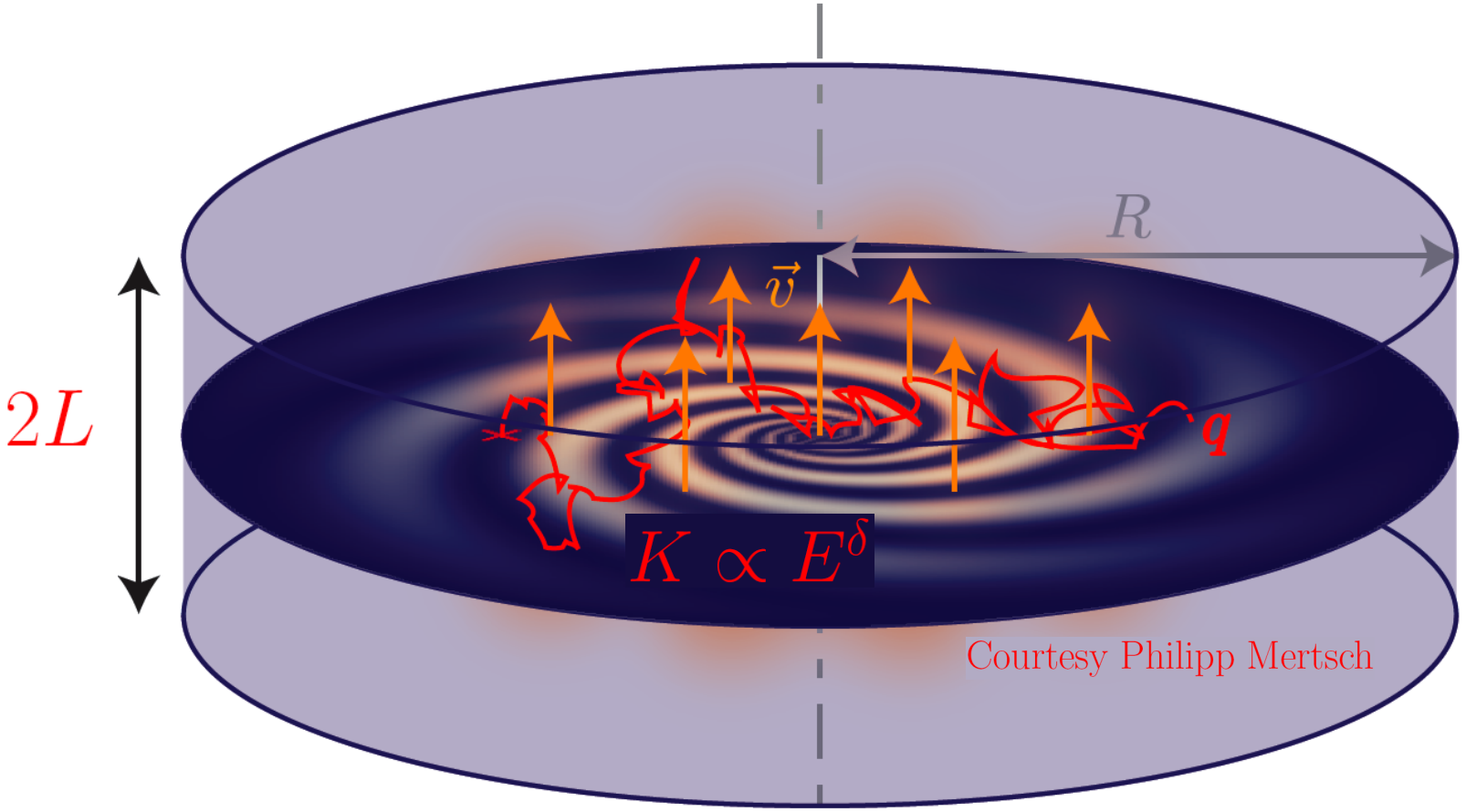
Backgrounds to a putative DM signal need to understood

Production cross sections – solar modulation – cosmic ray propagation



2) A new estimate of the antiproton background

Milky–Way seen by a cosmic–ray physicist



Cosmic rays propagate inside a diffusive halo

D. Maurin, R. Taillet, F. Donato, P. Salati, A. Barrau and G. Boudoul, *Galactic cosmic ray nuclei as a tool for astroparticle physics*, [astro-ph/0212111].

Antiprotons Production in the Galaxy

• **Secondary** antiprotons are produced through the spallations of cosmic–ray protons and He nuclei on the interstellar material.

$$p\,(\mathrm{CR}) \;+\; H\,(\mathrm{ISM}) \;\rightarrow\; \bar{p} \;+\; X$$

$$\Downarrow$$

$$q^{\mathrm{sec}}_{\bar{\mathrm{p}}}(r, E_{\bar{\mathrm{p}}}) \;=\; 4\pi \int_{E^0_{\mathrm{p}}}^{+\infty} \frac{d\sigma_{\mathrm{pH}\rightarrow\bar{\mathrm{p}}}}{dE_{\bar{\mathrm{p}}}} \{E_{\mathrm{p}} \rightarrow E_{\bar{\mathrm{p}}}\}\; n_{\mathrm{H}}\, \Phi_{\mathrm{p}}(r, E_{\mathrm{p}})\, dE_{\mathrm{p}}$$

New developments since 2008

– CR p and He fluxes measured with improved accuracy.

$$\Phi^{\mathrm{sec}}_{\bar{\mathrm{p}}}(E_{\bar{\mathrm{p}}}) \;\propto\; q^{\mathrm{sec}}_{\bar{\mathrm{p}}}(E_{\bar{\mathrm{p}}}) \;\propto\; \Phi_{\mathrm{p}}(E_{\mathrm{p}} \sim 10 \times E_{\bar{\mathrm{p}}})$$

$$\Downarrow$$

$$\bar{\mathrm{p}}/\mathrm{p}(E) \;\propto\; \Phi_{\mathrm{p}}(\sim 10 \times E)/\Phi_{\mathrm{p}}(E) \;\propto\; 10^{-\alpha}$$

$\bar{\mathrm{p}}/\mathrm{p}$ depends on the CR proton spectral index α

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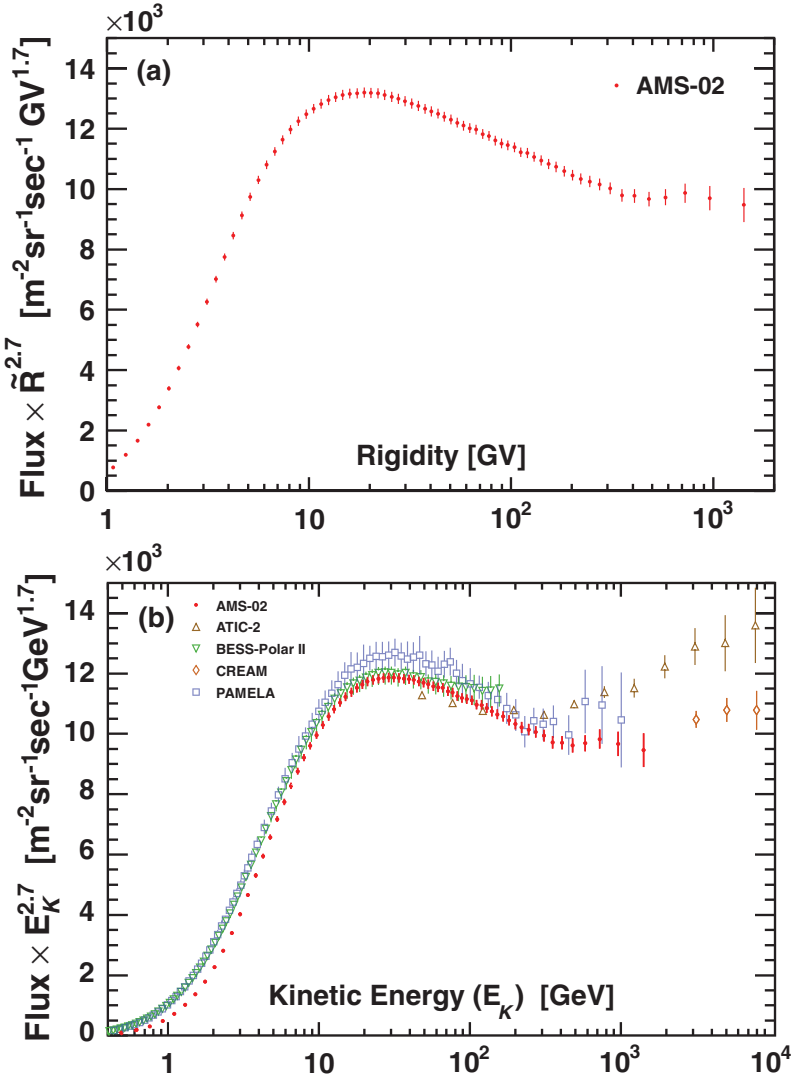
Precision Measurement of the Proton Flux in Primary Cosmic Rays from Rigidity 1 GV to 1.8 TV with the Alpha Magnetic Spectrometer on the International Space Station



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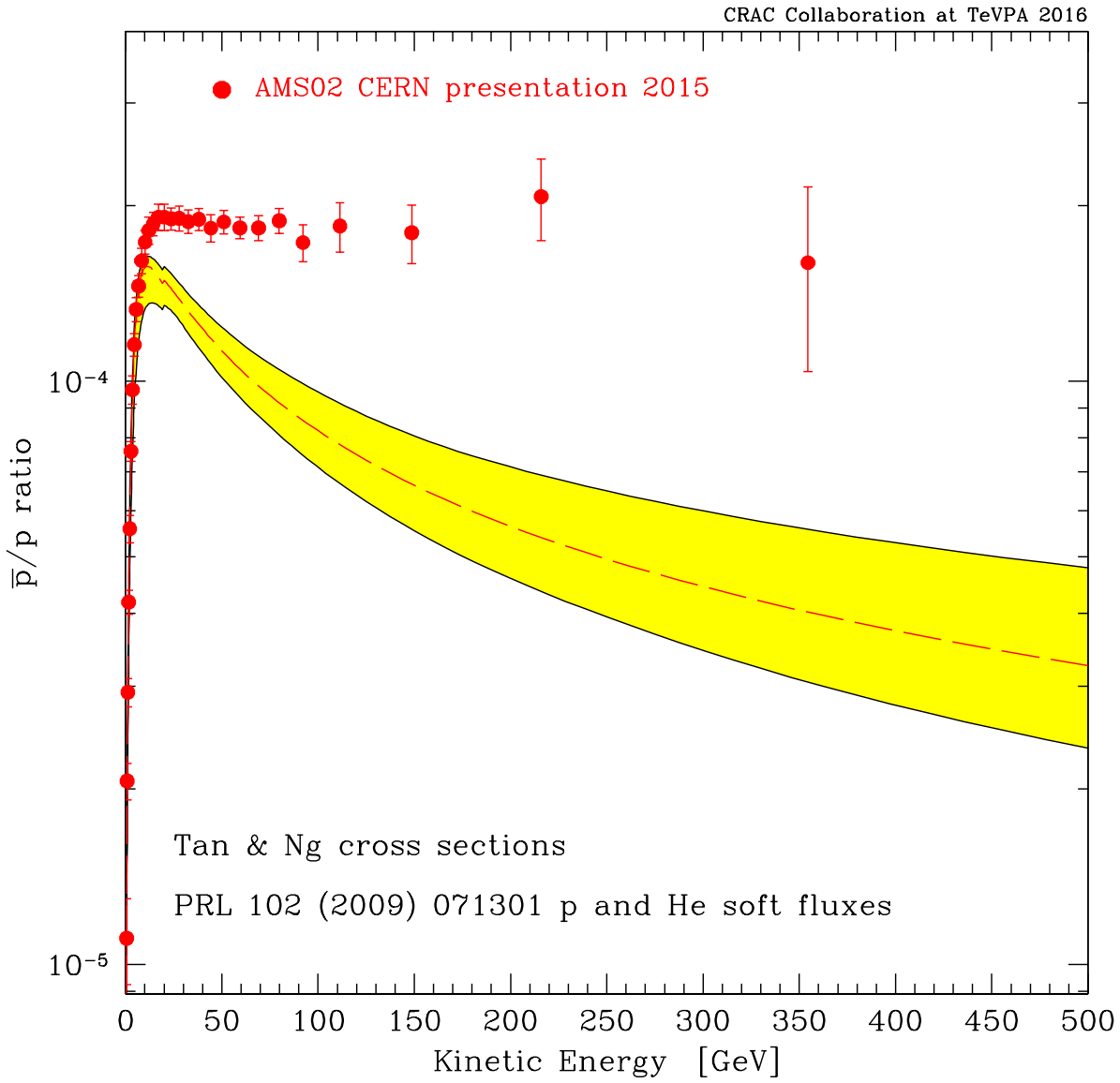

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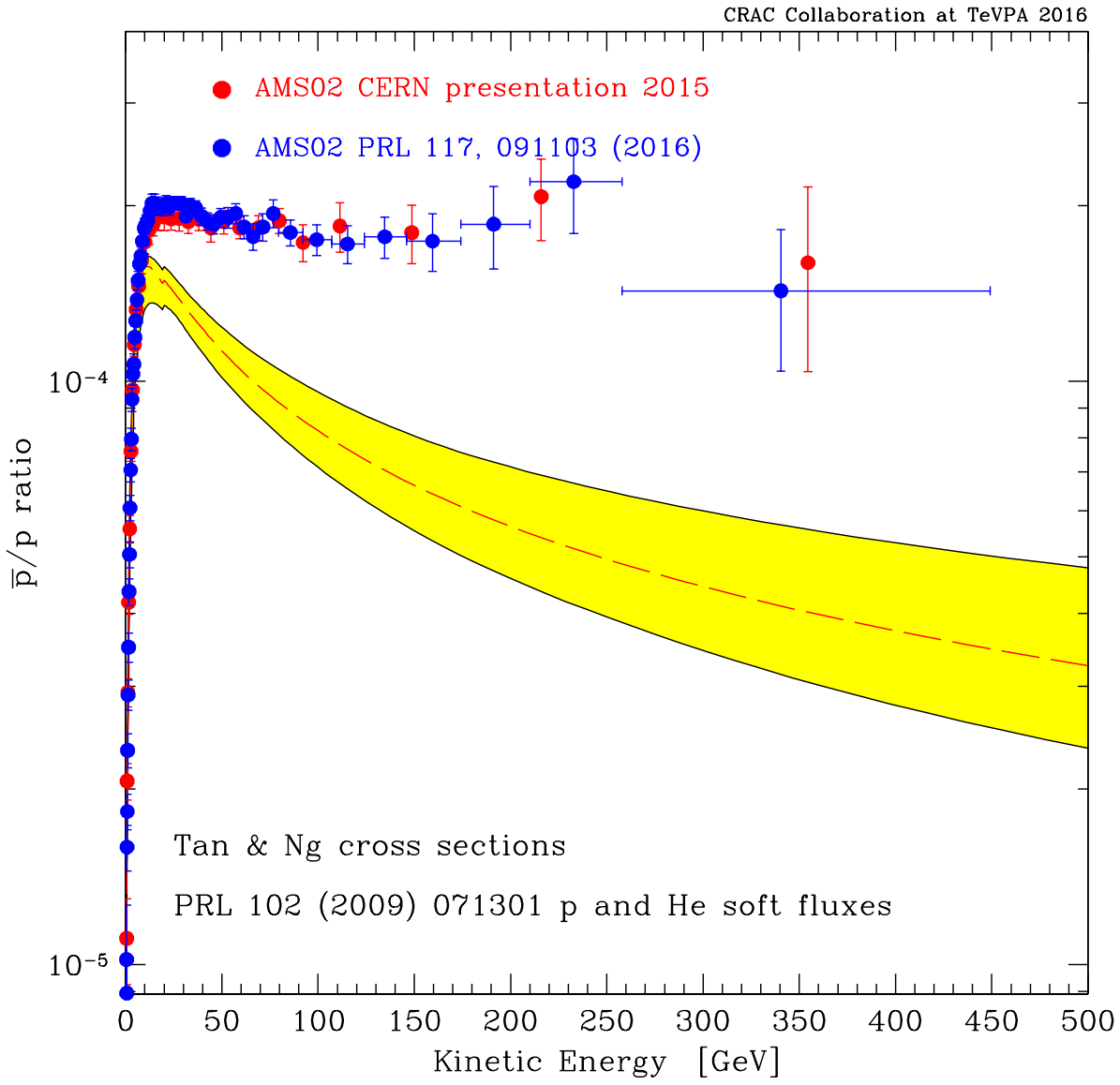

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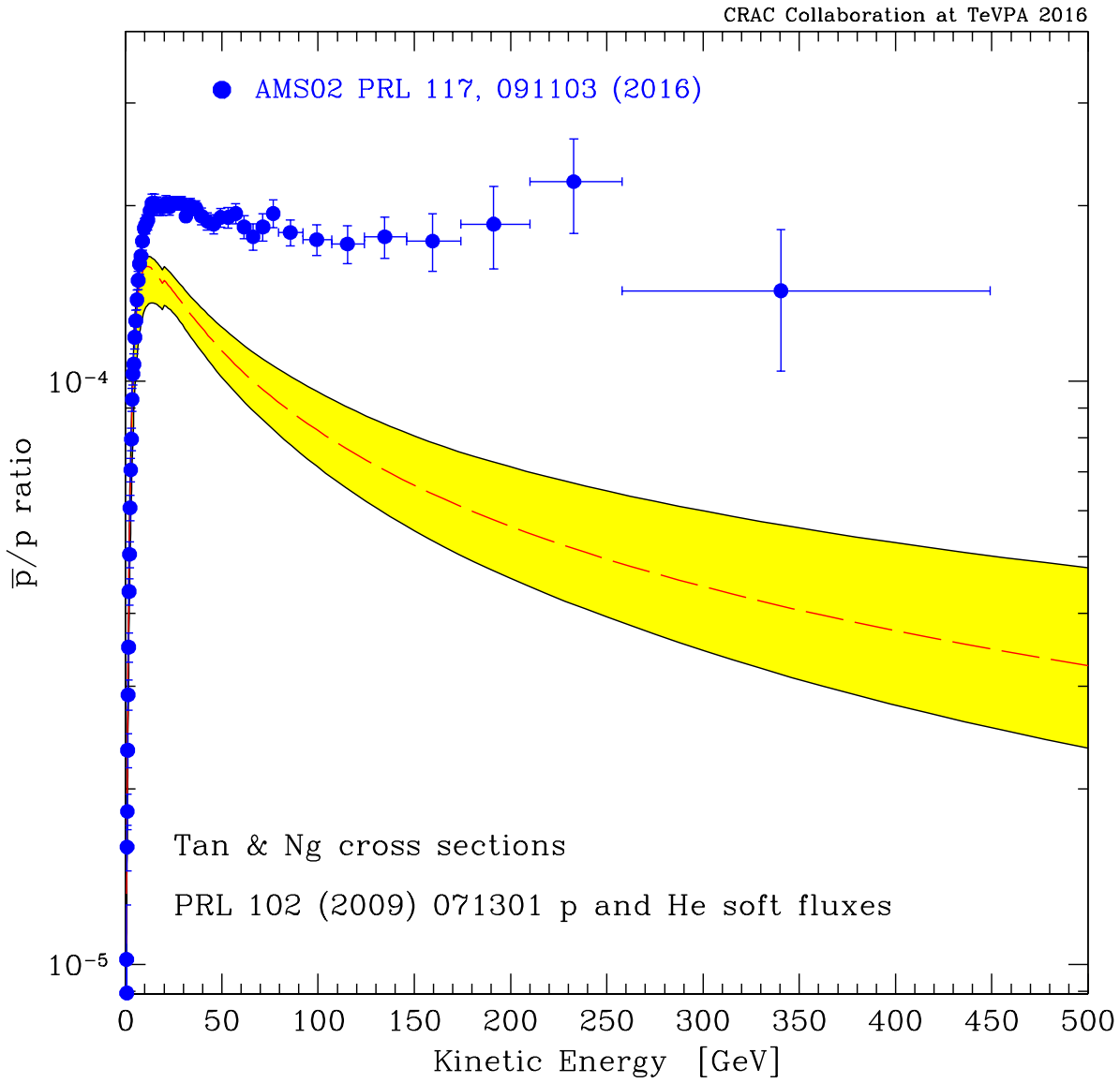

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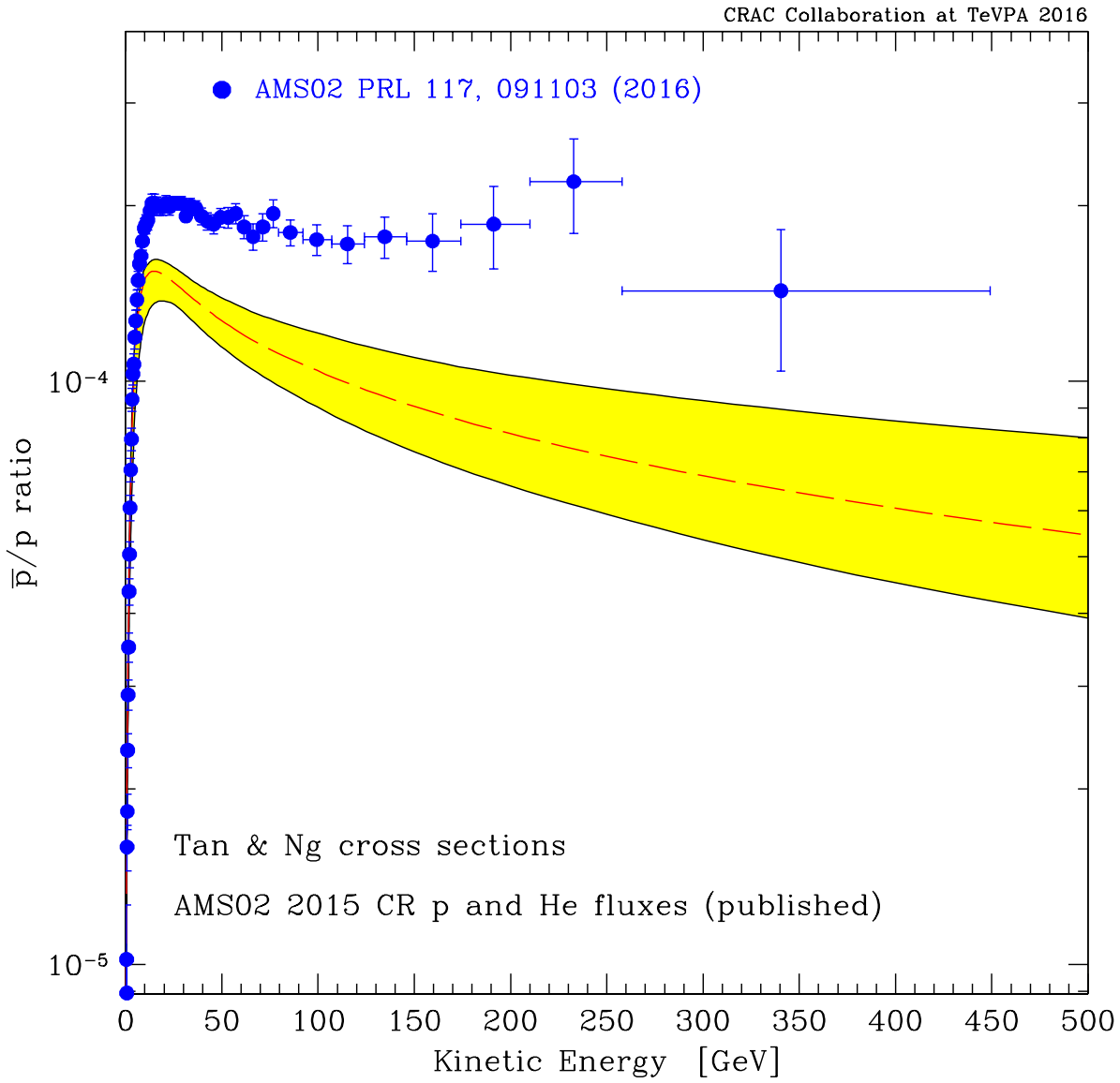

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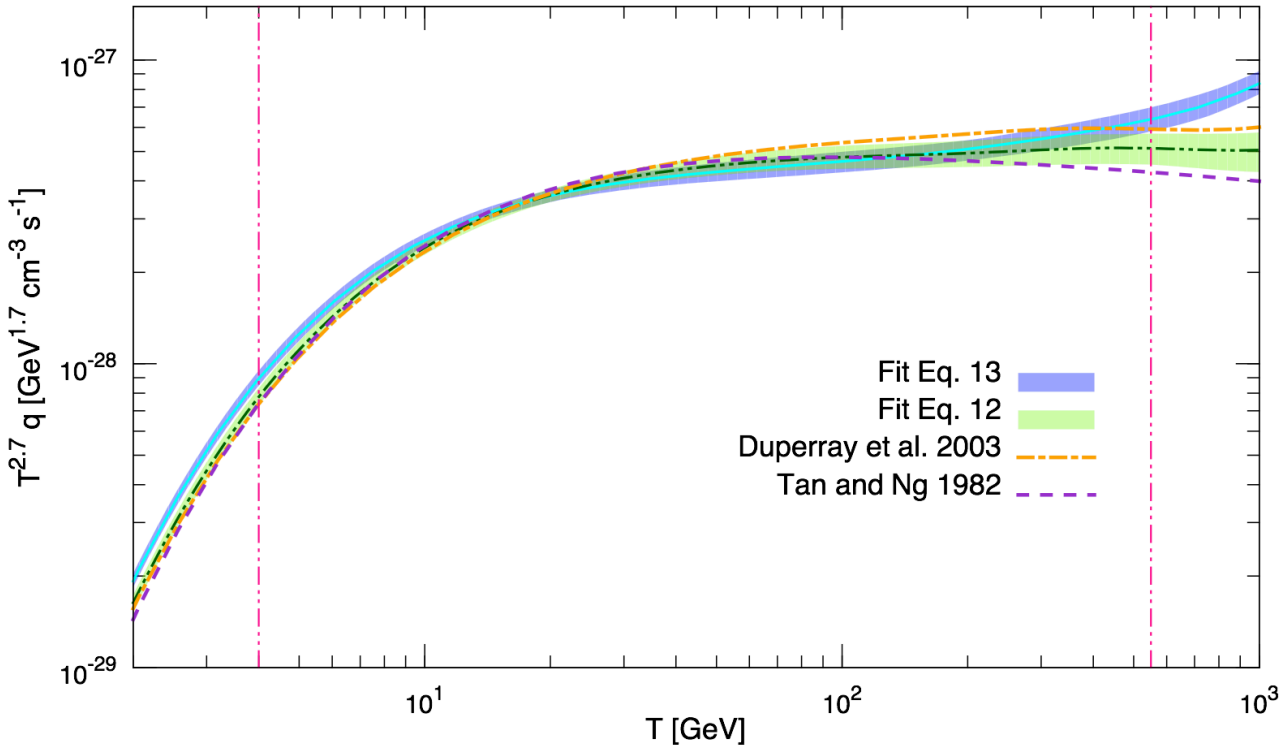
– New parameterization of $d\sigma_{\mathrm{pH}\rightarrow\bar{\mathrm{p}}}/dE_{\bar{\mathrm{p}}}$ from BRAHMS and NA49.



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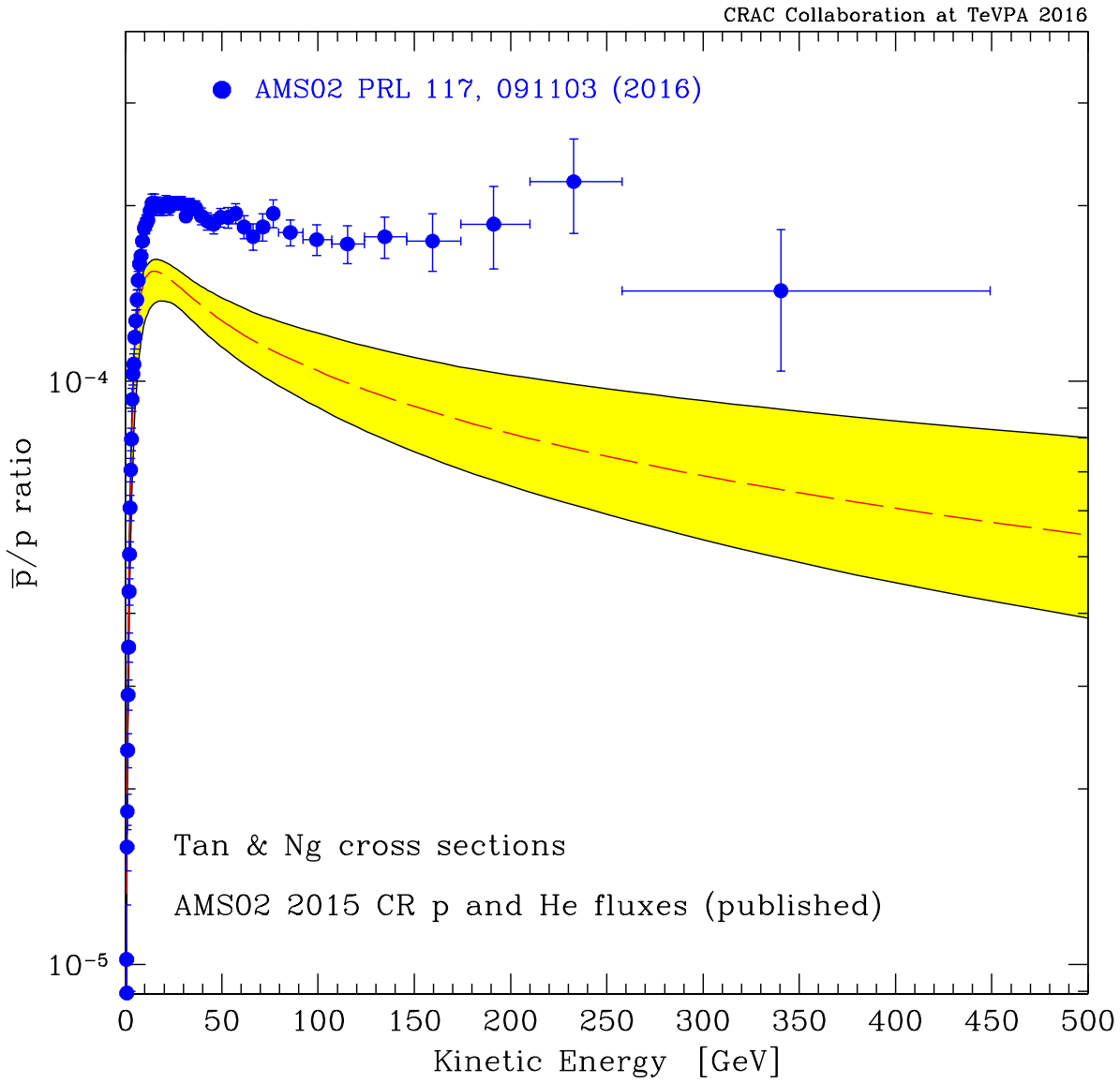
A new evaluation of the antiproton production cross section for cosmic ray studies

Mattia di Mauro[1,2], Fiorenza Donato[1], Andeas Goudelis[2], Pasquale Dario Serpico[2]

$$E\frac{d^3\sigma}{dp^3} = \sigma_{\mathrm{in}}(s)(1-x_R)^{C_1}e^{-C_2x_R} \qquad (12)$$
$$\left[C_3(\sqrt{s})^{C_4}e^{-C_5p_T} + C_6(\sqrt{s})^{C_7}e^{-C_8p_T^2}\right].$$

C_1 (error)	C_2 (error)	C_3 (error)	C_4 (error)	C_5 (error)	C_6 (error)	C_7 (error)	C_8 (error)
4.499(0.040)	3.41(0.11)	0.00942(0.00083)	0.445(0.027)	3.502(0.018)	0.0622(0.0086)	-0.247(0.049)	2.576(0.027)

New developments since 2008



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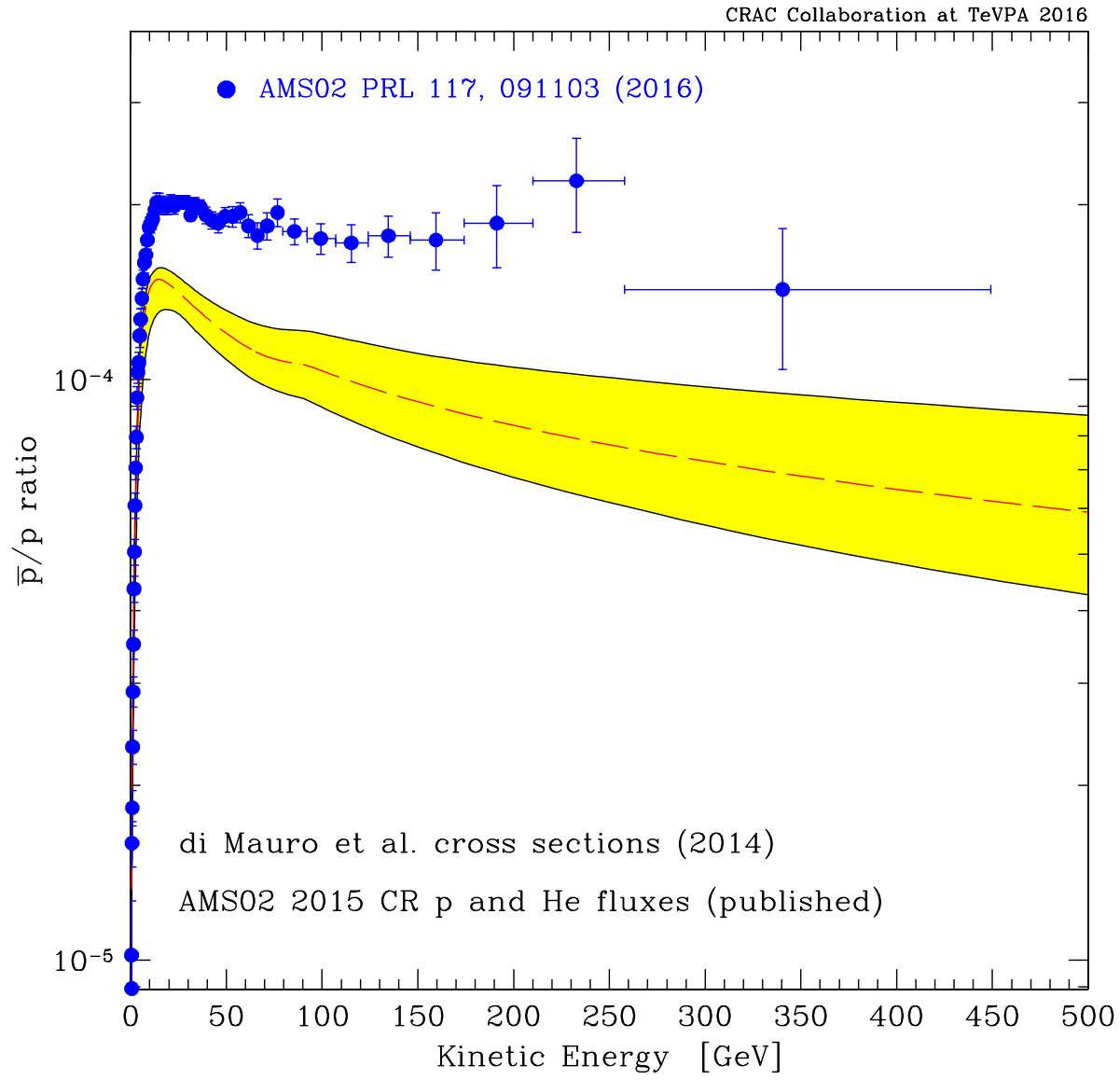

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New developments since 2008

– Proton collisions yield more antineutrons than antiprotons.

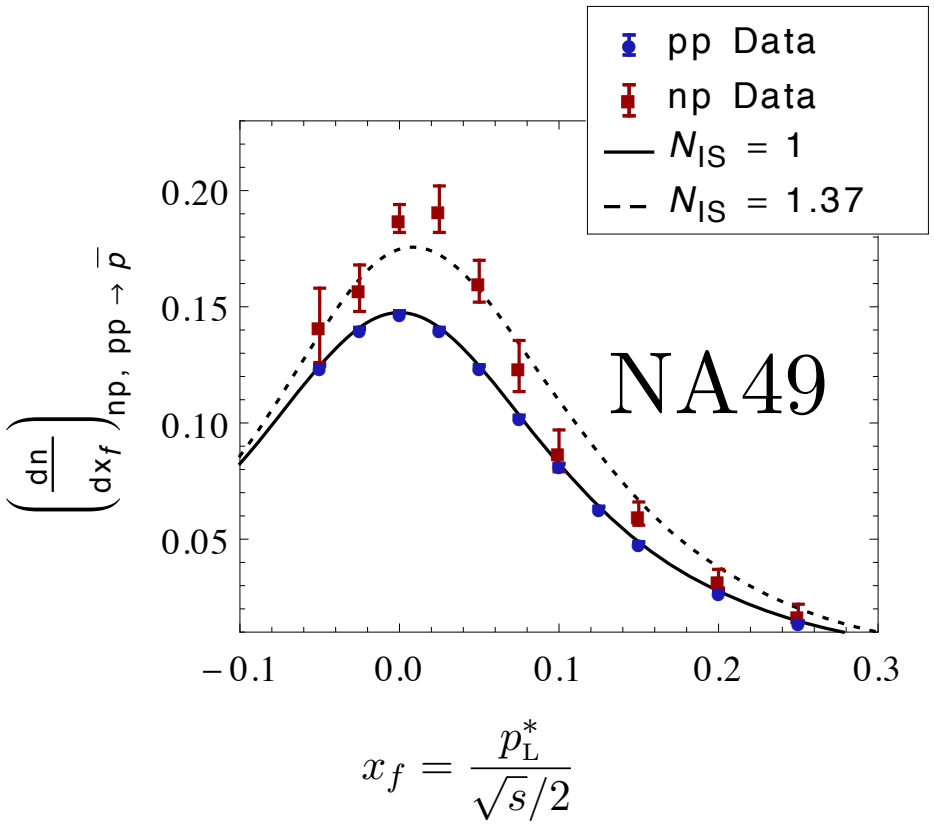
$$d\sigma_{\mathrm{pp}\rightarrow\bar{\mathrm{n}}}/dE_{\bar{\mathrm{p}}} \equiv d\sigma_{\mathrm{pn}\rightarrow\bar{\mathrm{p}}}/dE_{\bar{\mathrm{p}}}$$

$$\Downarrow$$

$$d\sigma_{\mathrm{pp}\rightarrow\bar{\mathrm{n}}}/d\sigma_{\mathrm{pp}\rightarrow\bar{\mathrm{p}}} \simeq 1.3$$

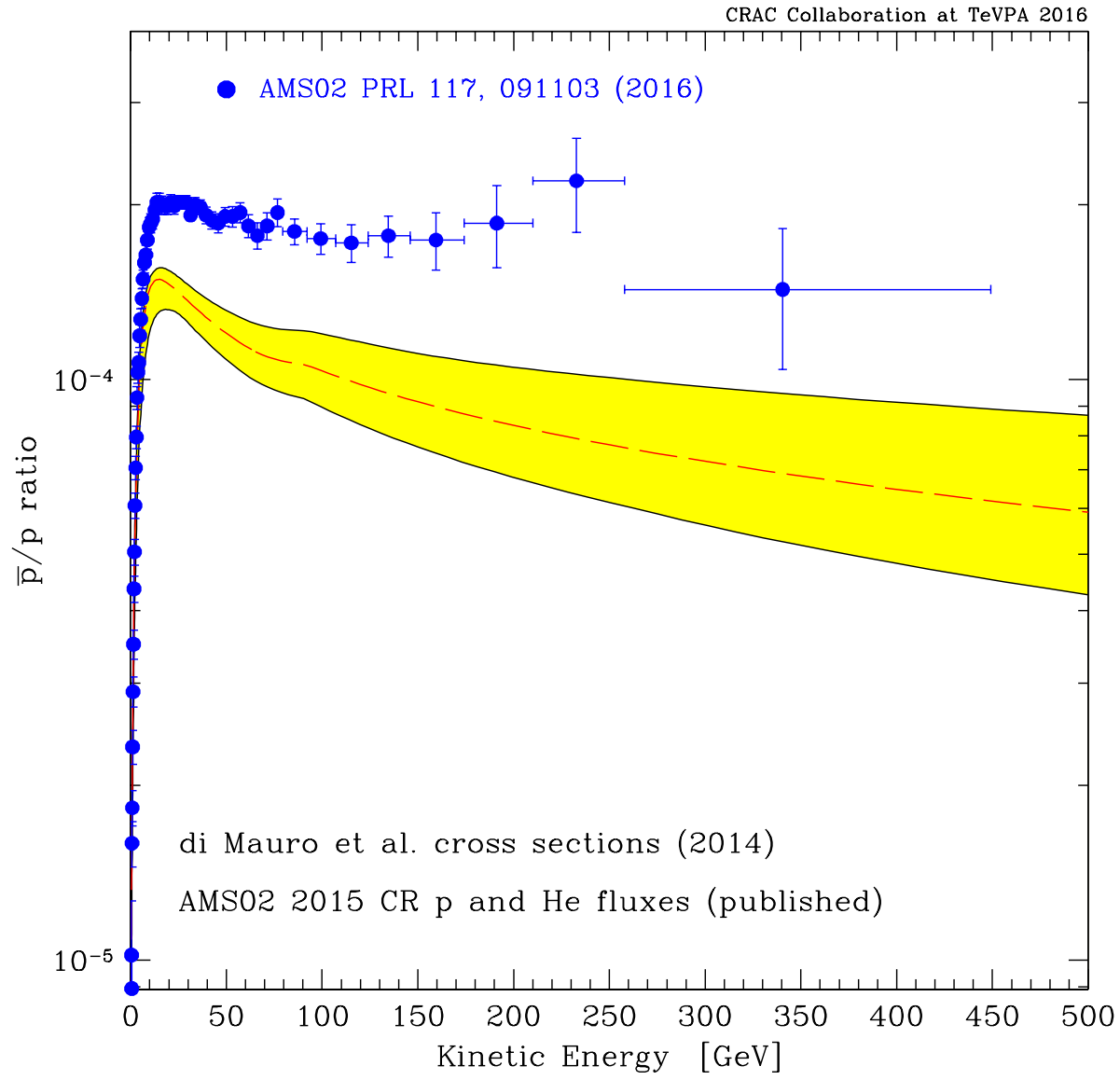
The Cosmic Ray Antiproton Background for AMS-02

Rolf Kappl[a], **Martin Wolfgang Winkler**[b]



NA49 Collaboration, H. Fischer, Heavy Ion Phys. **17** (2003), 369–386

New developments since 2008



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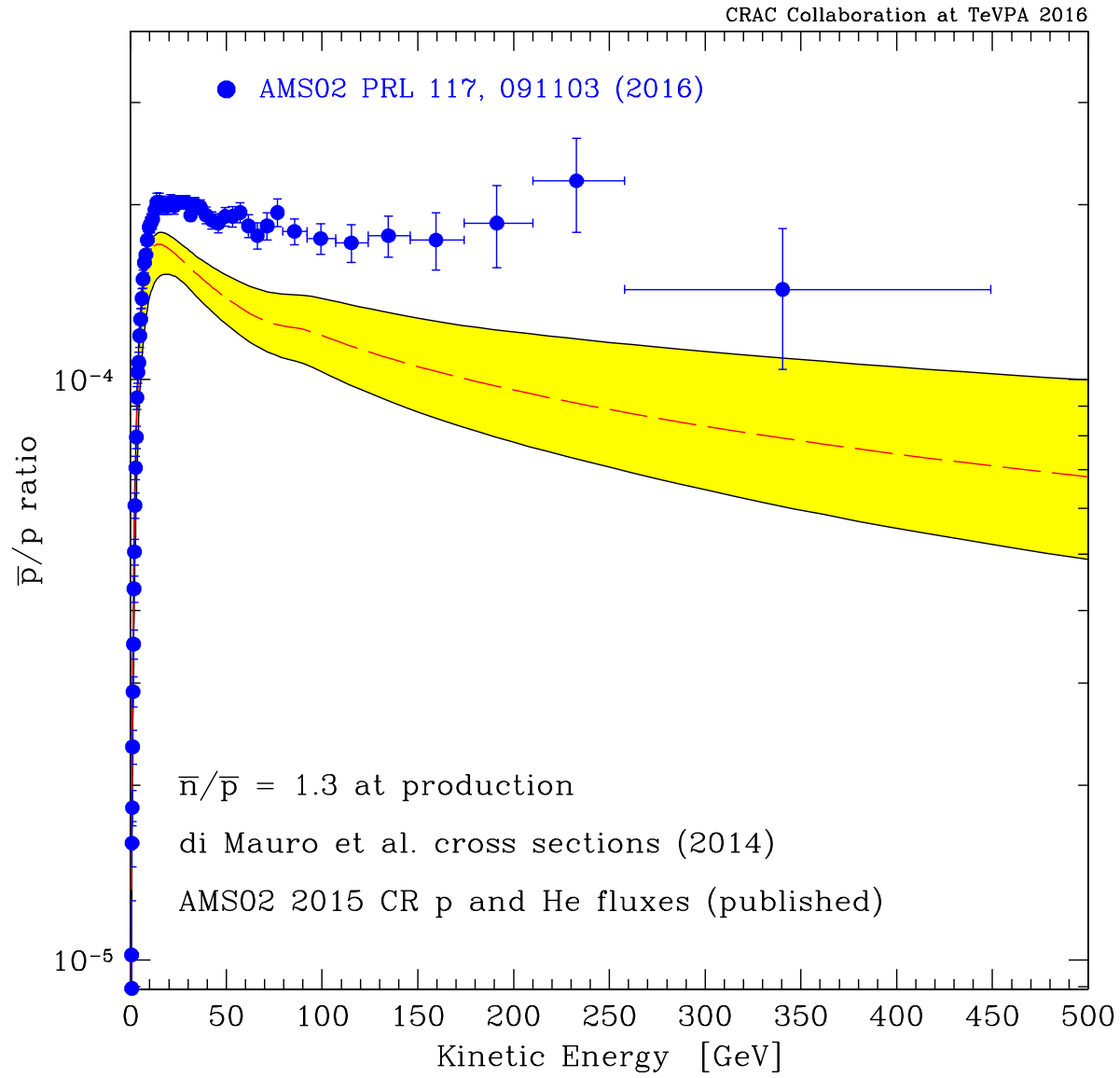

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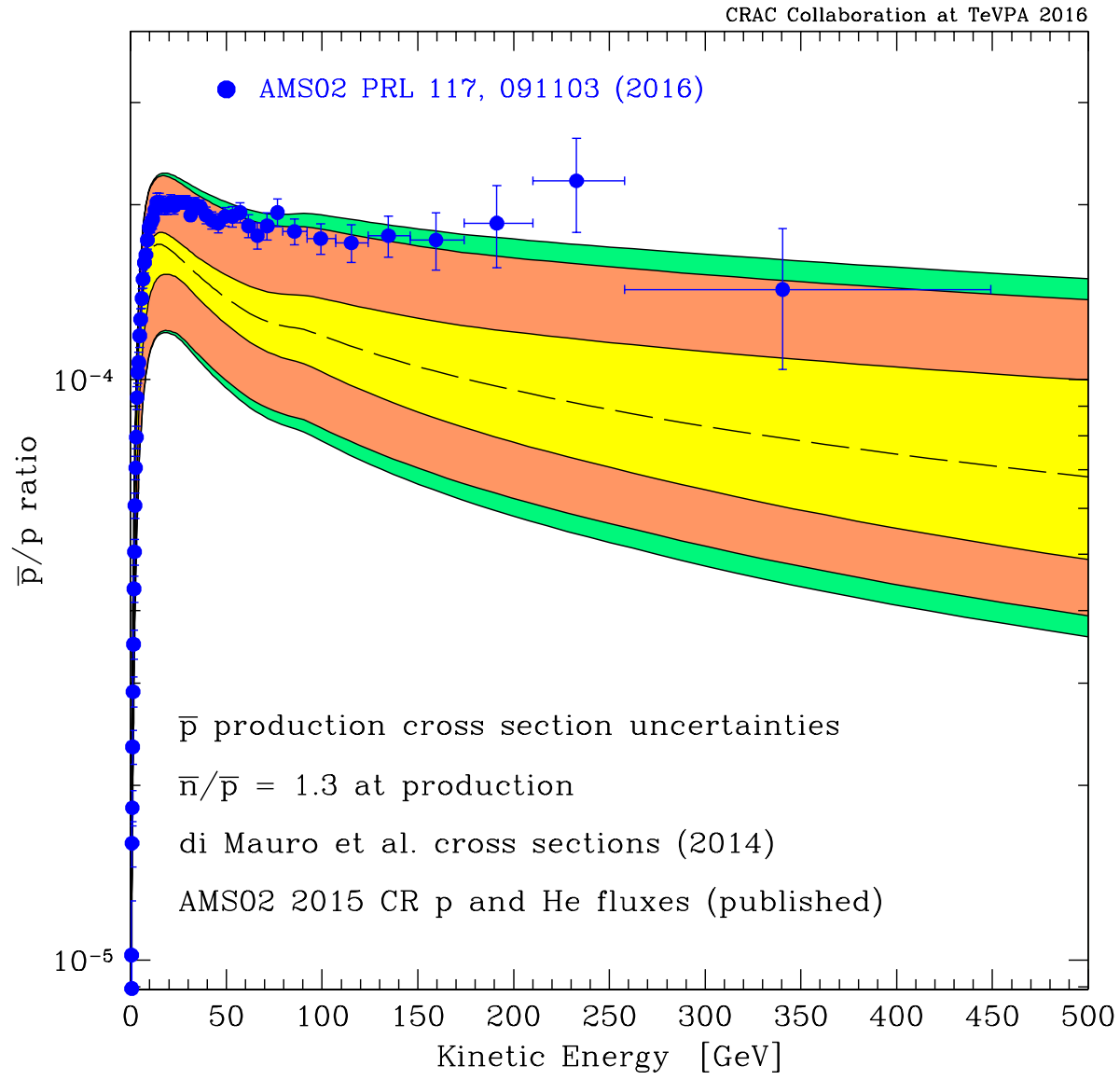
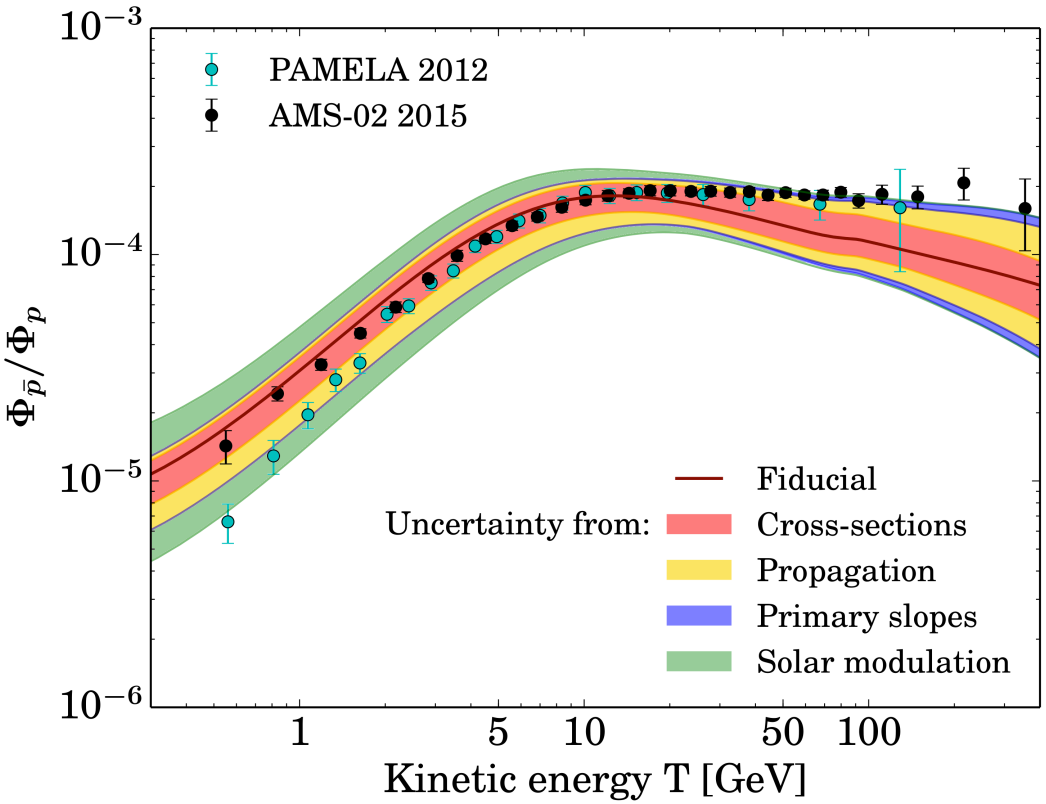


Antiproton data are compatible with the background

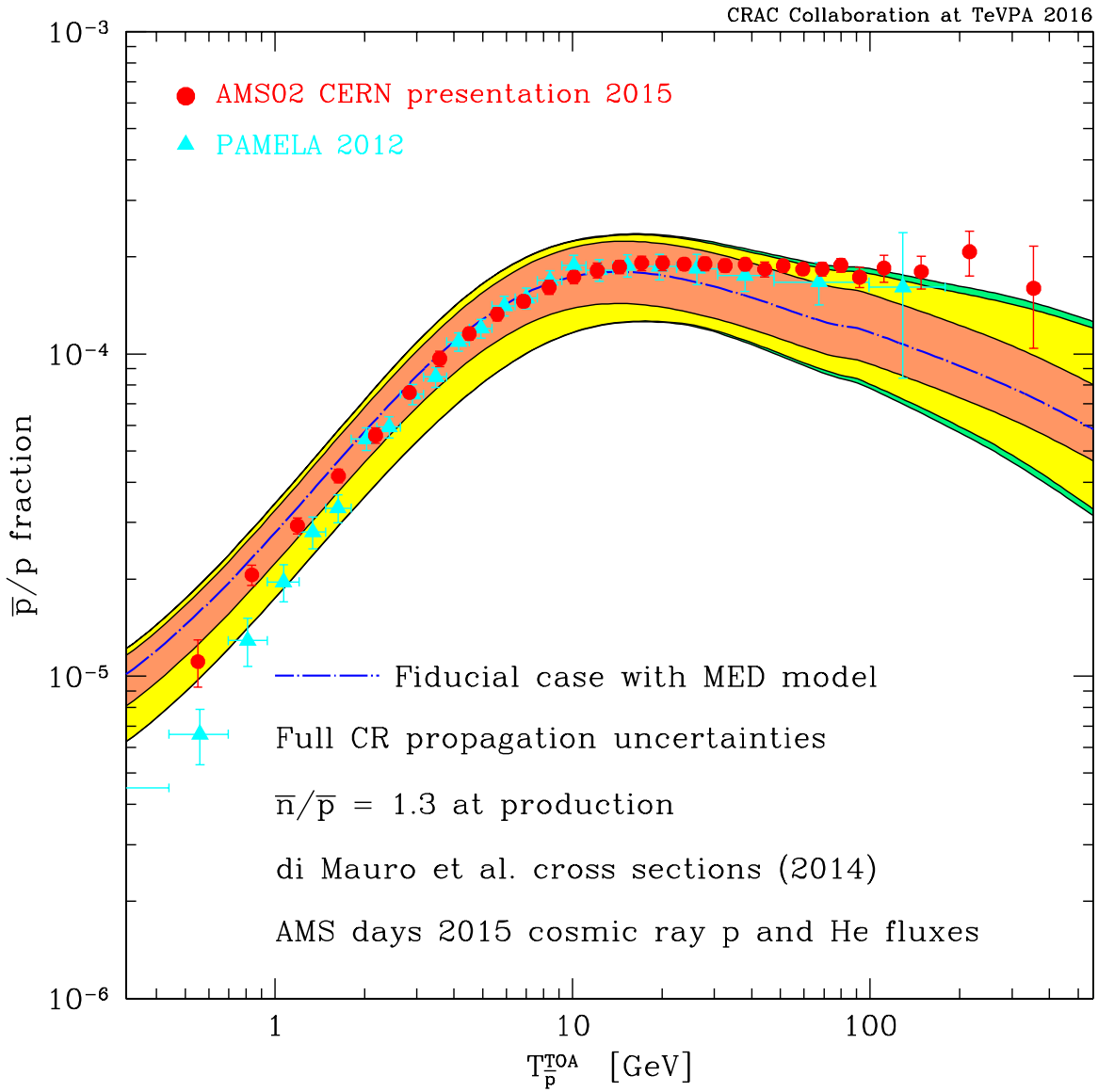
No $\bar{p}$ excess



But measurements are on the upper side

arXiv:1504.04276v2 [astro-ph.HE] 28 Apr 2015

Recent ongoing developments with published data



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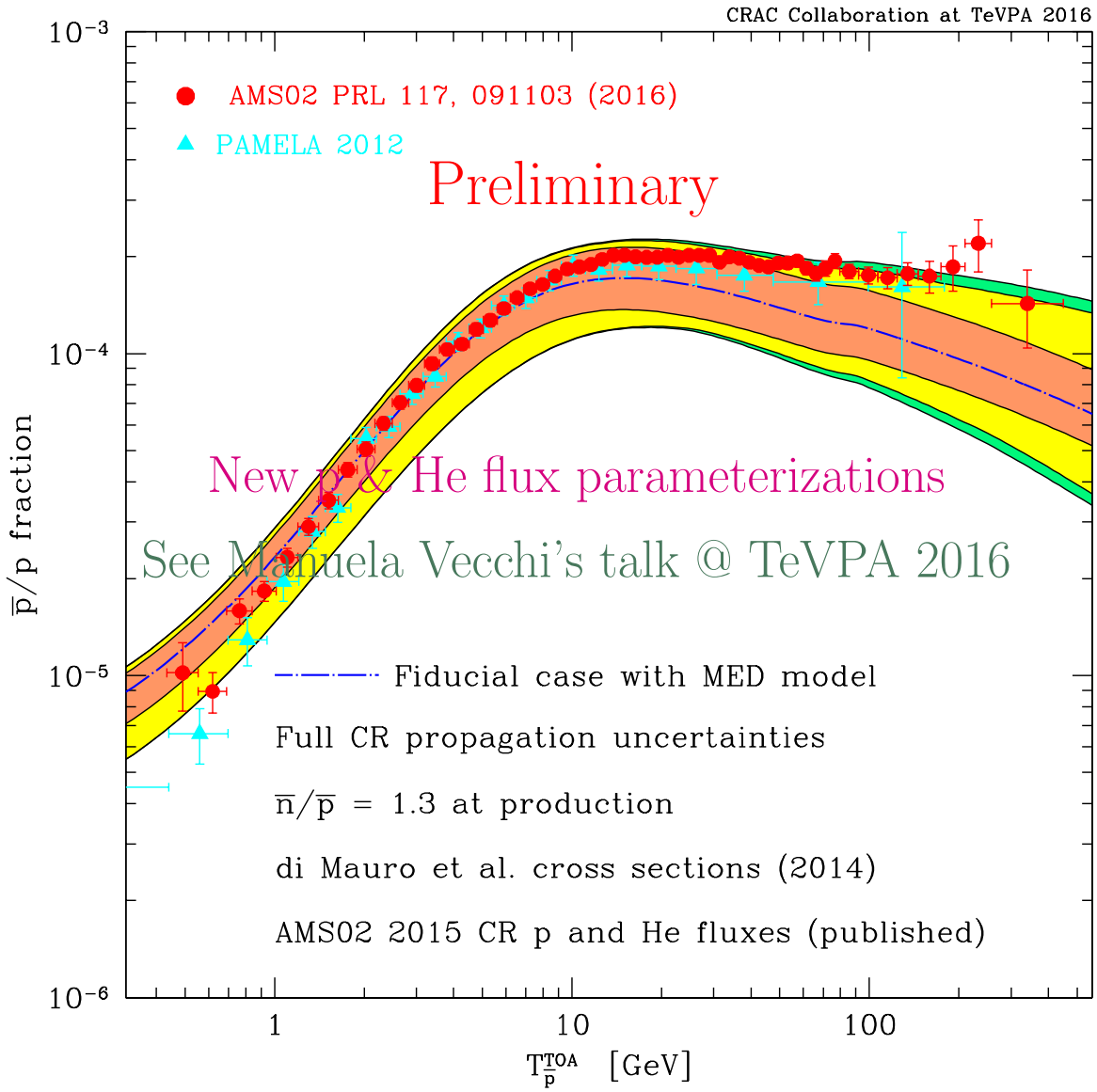

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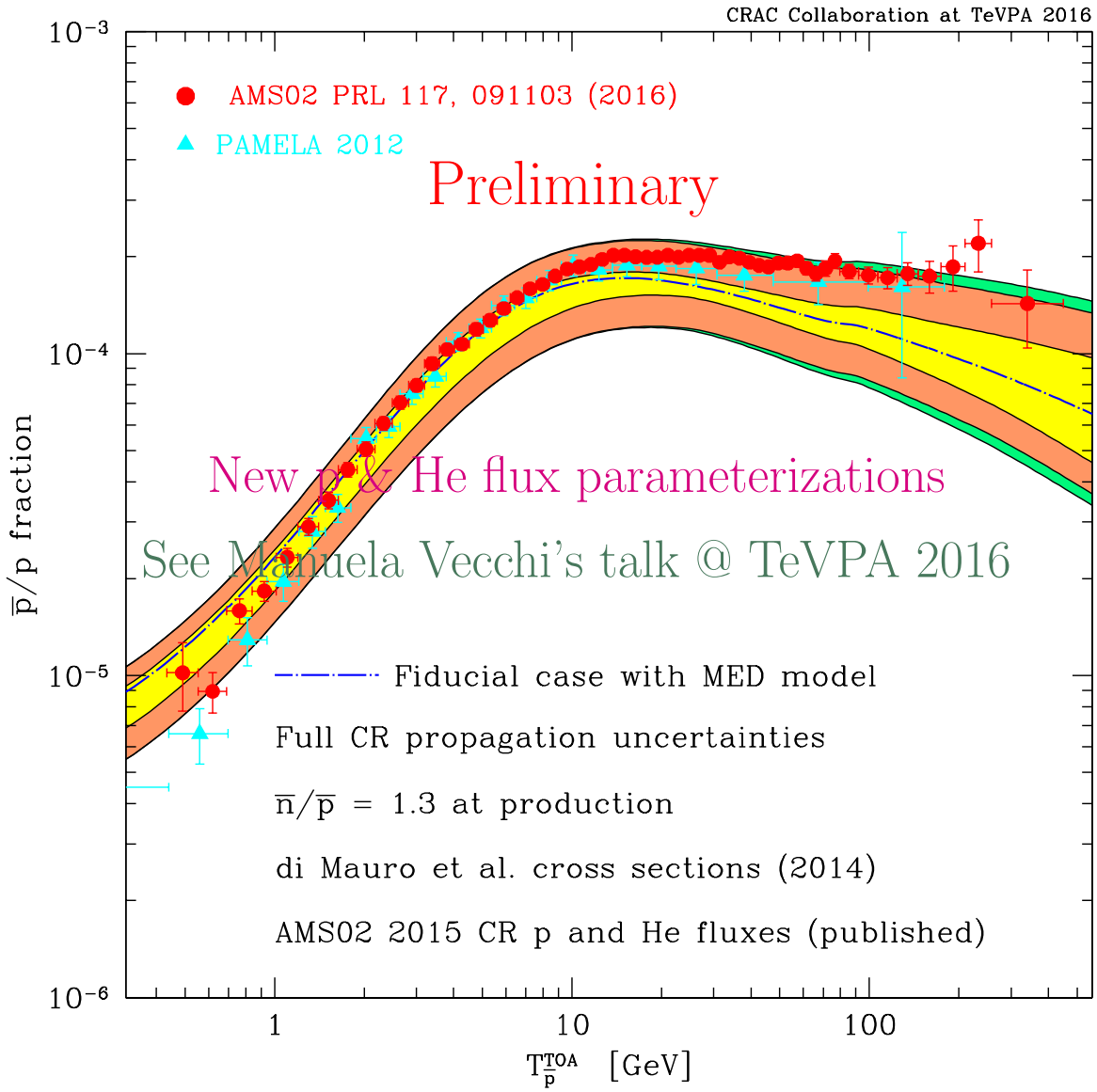
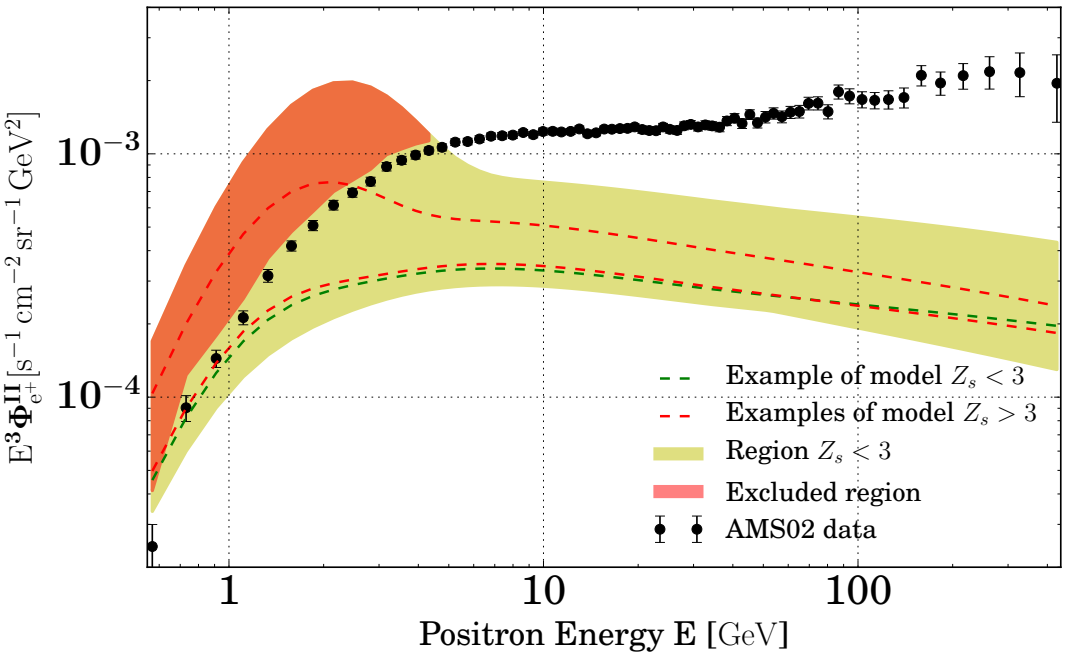


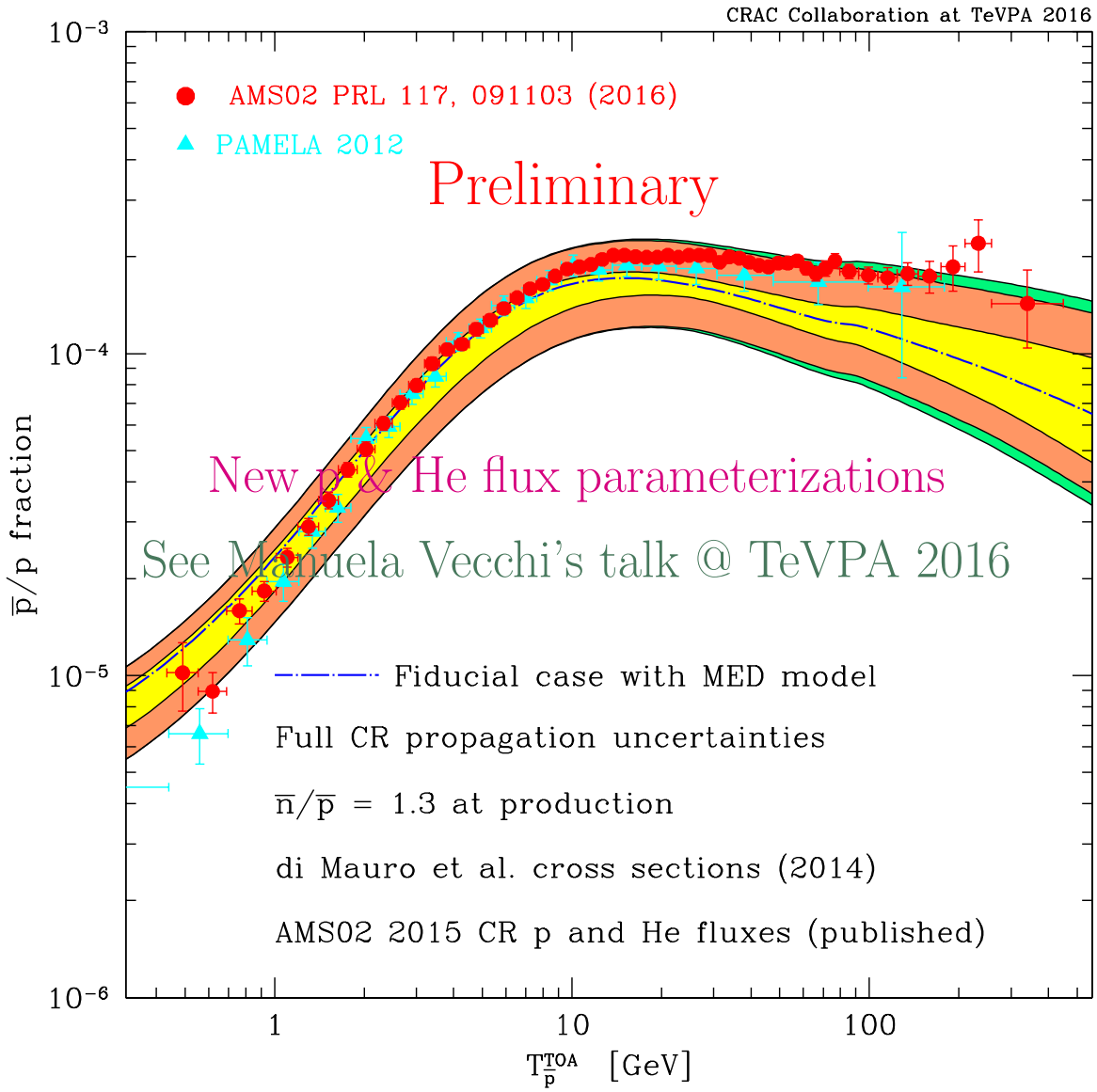
Preliminary



Positron flux constrains CR propagation

See Mathieu Boudaud's talk @ TeVPA 2016

Recent ongoing developments with published data



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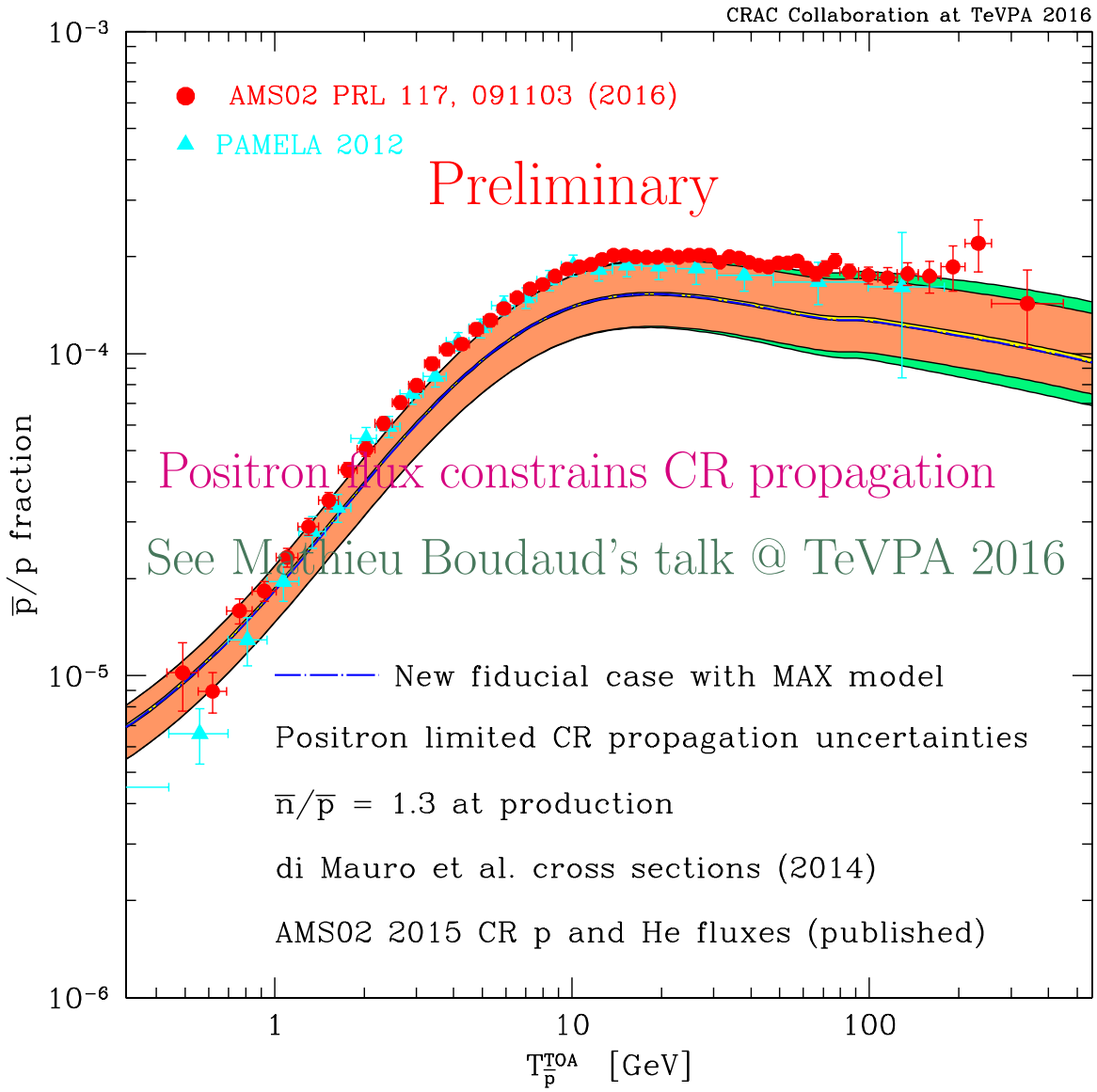

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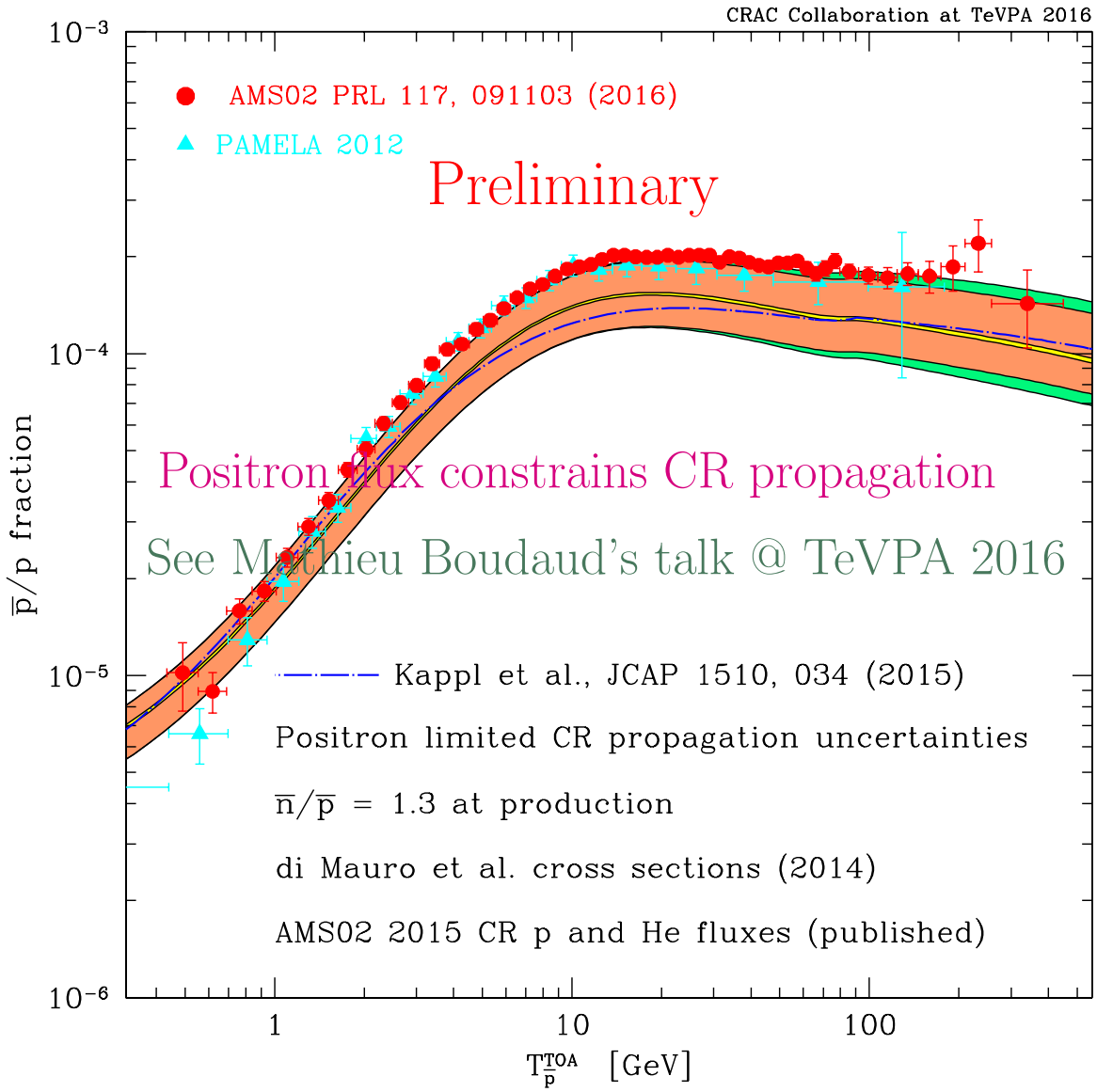

3) Limits on DM properties – a status report

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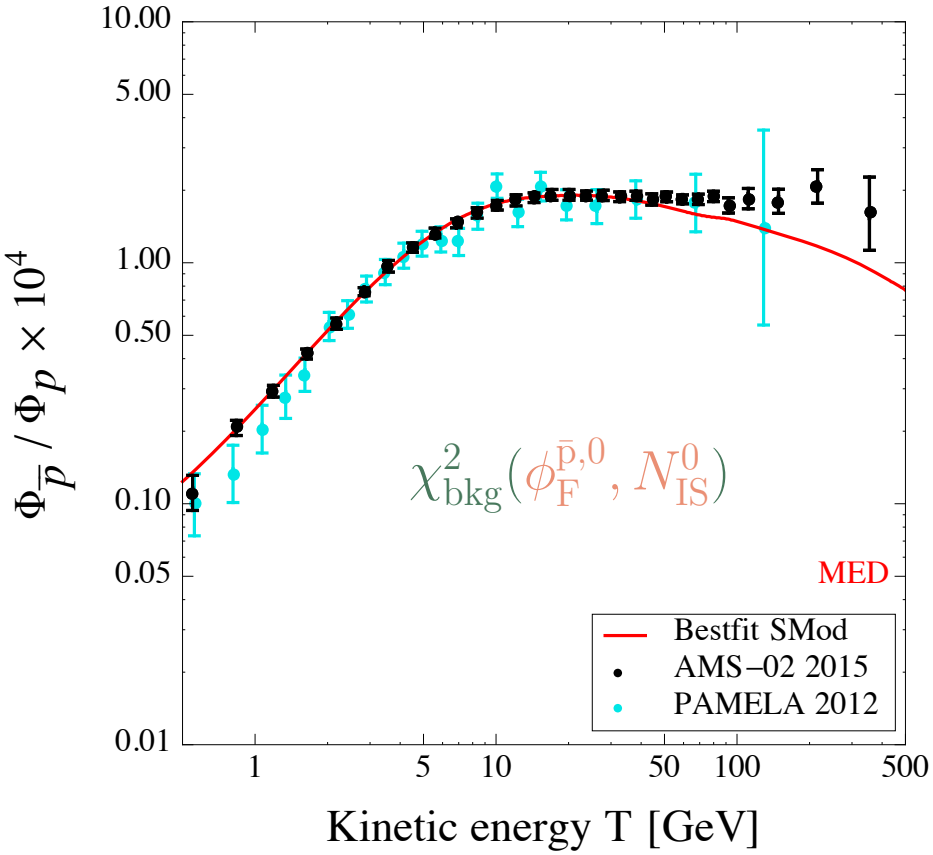
• **Primary** antiprotons originate from the annihilations of the dark matter species – WIMPs in our case – concealed in the Galactic halo.

$$\chi \;+\; \chi \;\rightarrow\; q\,\bar{q},\, W^+W^-,\, HH, \ldots \;\rightarrow\; \bar{p} \;+\; X$$

$$\Downarrow$$

$$q^{\mathrm{susy}}_{\bar{\mathrm{p}}}(r, z, E_{\bar{\mathrm{p}}}) \;=\; \frac{1}{2} \langle \sigma_{\mathrm{ann}} v \rangle\, g(T_{\bar{\mathrm{p}}}) \left\{ \frac{\rho_\chi(r, z)}{m_\chi} \right\}^2$$

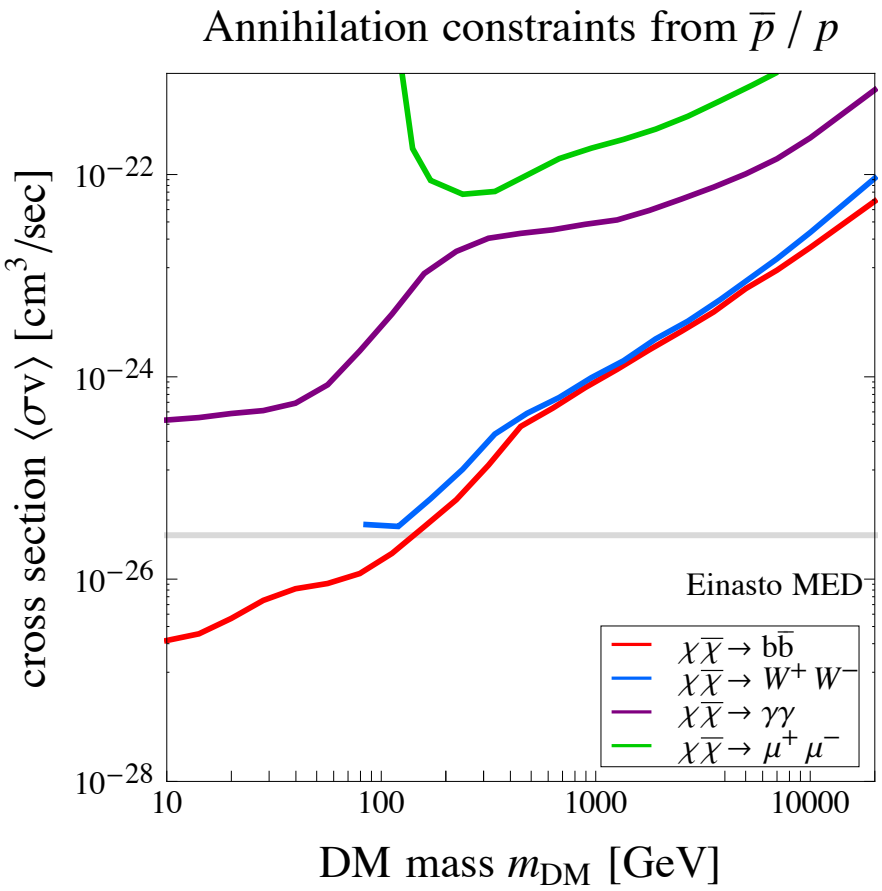
Case	δ	K_0 [kpc^2/Myr]	L [kpc]	V_C [km/s]	V_a [km/s]
max	0.46	0.0765	15	5	117.6
med	0.70	0.0112	4	12	52.9
min	0.85	0.0016	1	13.5	22.4



New limits can be set on DM properties

Treat $\phi_{\rm F}^{\bar{\rm p}}$ and $N_{\rm IS} = \bar{\rm n}/\bar{\rm p}$ as nuisance parameters

$$\chi^2(m_{\rm DM}, \langle\sigma v\rangle, \phi_{\rm F}^{\bar{\rm p}}, N_{\rm IS}) - \chi^2_{\rm bkg}(\phi_{\rm F}^{\bar{\rm p},0}, N_{\rm IS}^0) \leq 4$$



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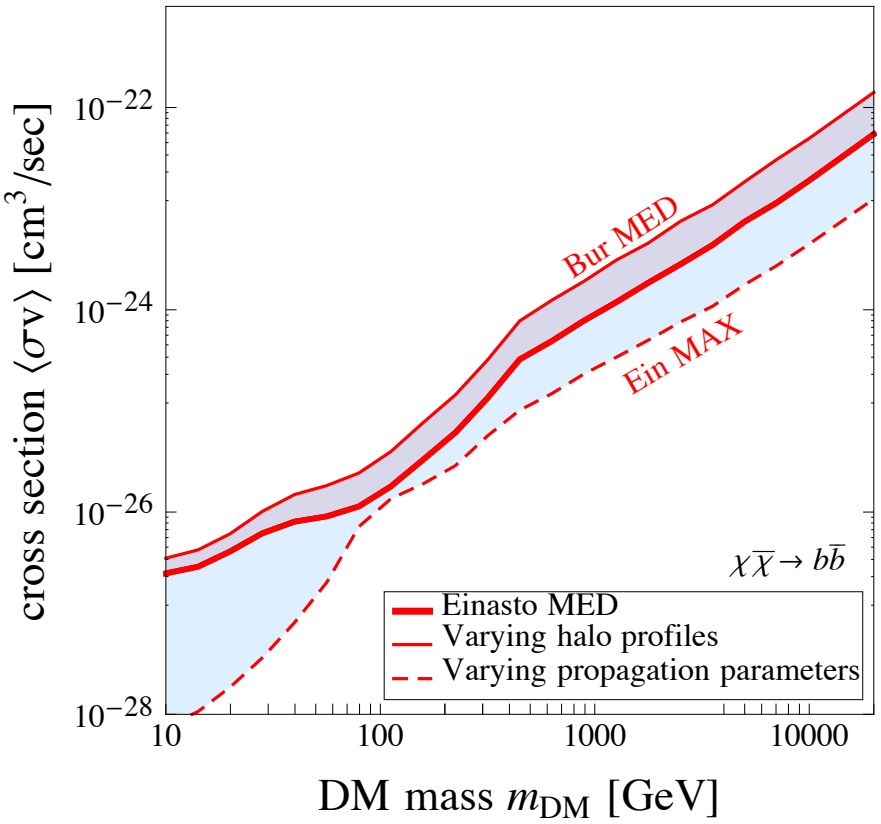
Astrophysical uncertainties on the constraints



New limits can be set on DM properties

Closing thoughts

• Dark Matter indirect detection is a powerful probe provided that **astrophysical backgrounds** are well determined.

$$\Downarrow$$

Cosmic ray propagation & X-sections

• The antiproton background has been increasing somewhat since 2008...

(i) harder CR proton and He spectra,

(ii) increased $\bar{n}/\bar{p}$ ratio at production,

(iii) better determination of the cross-section uncertainties.

• The AMS-02 antiproton measurements are of exquisite quality. They are – at that stage – compatible with the background although at the upper limit of what is expected.

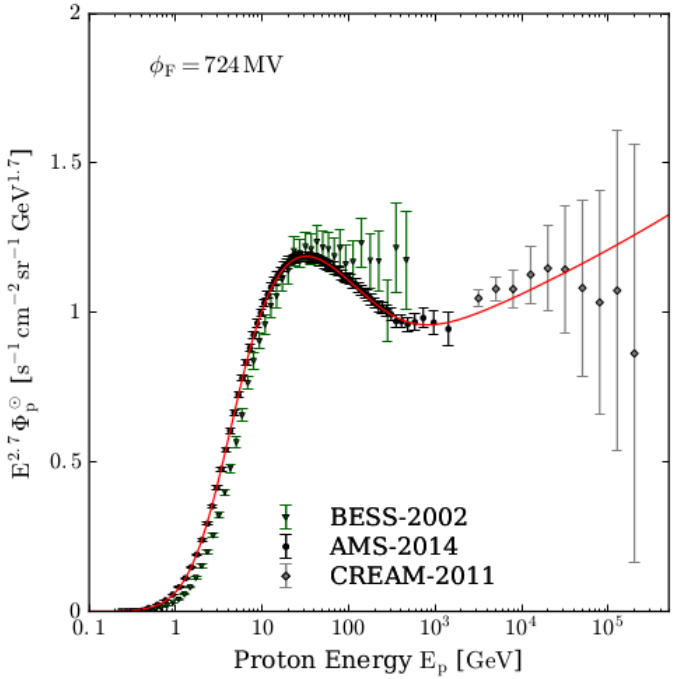
• To decide whether a signal is hidden inside the AMS-02 $\bar{p}/p$ data, CR propagation needs to be better constrained and the antiproton production cross-sections should be more accurately measured.

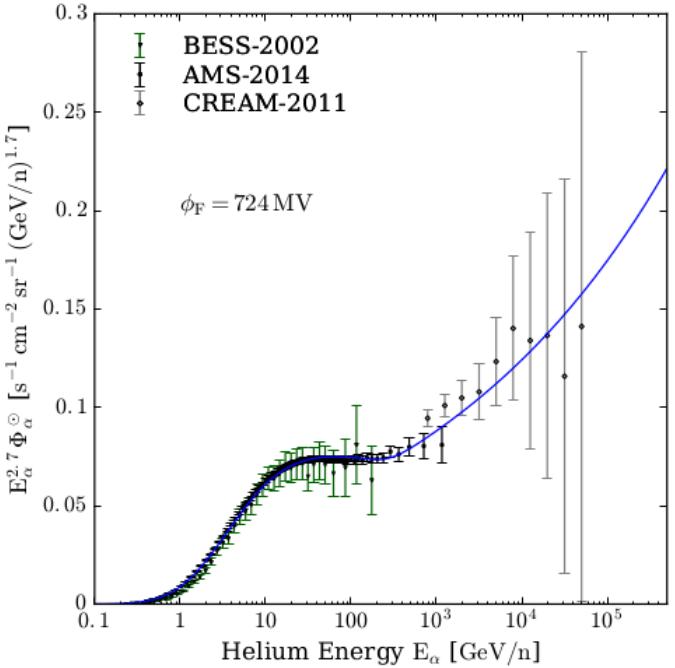
B/C (published) measurements are eagerly awaited

Back-up slides

Recent ongoing developments with published data

Preliminary





New p & He flux parameterizations

See Manuela Vecchi's talk @ TeVPA 2016

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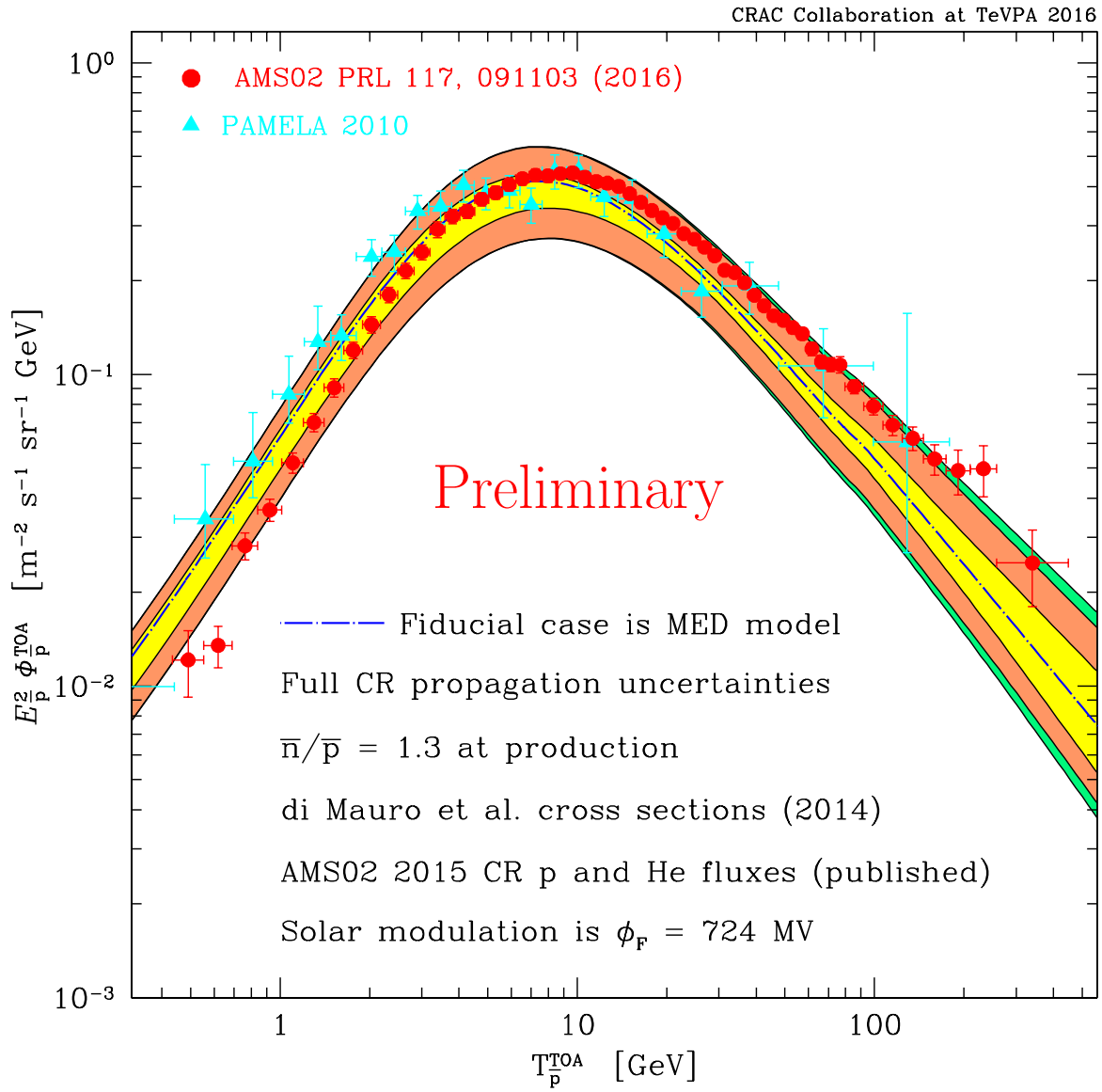

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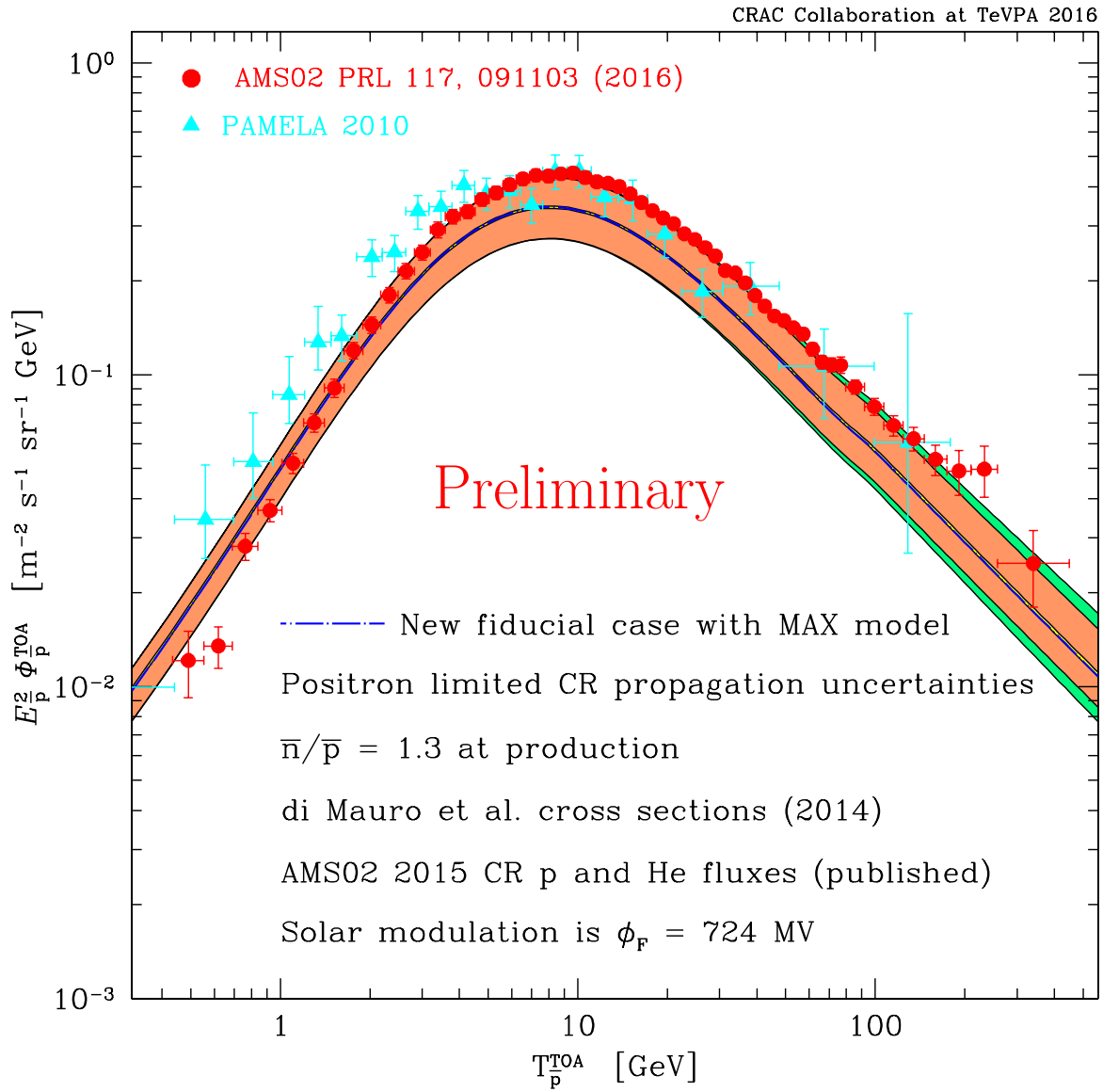

Recent ongoing developments with published data