2nd generation of gravitational wave interferometric detectors: the example of Advanced Virgo

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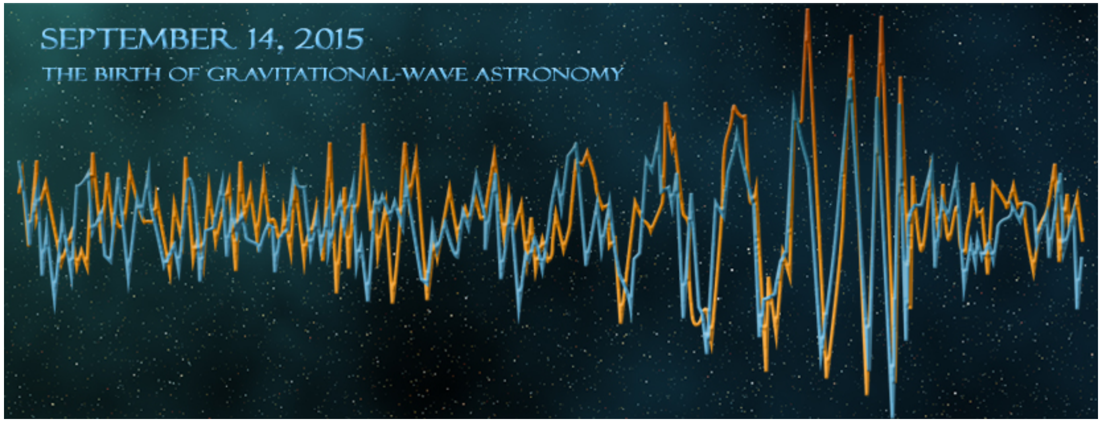
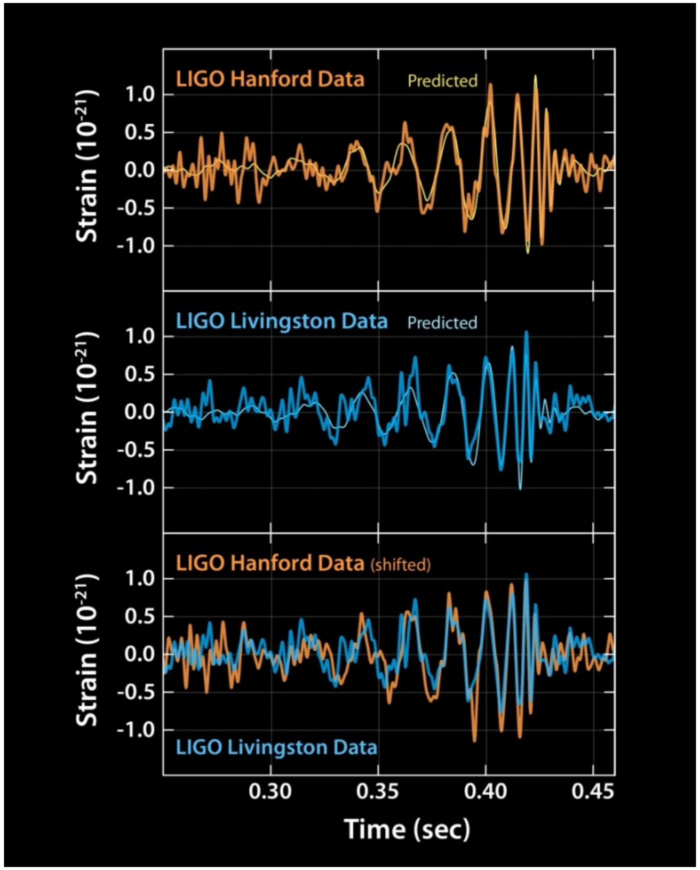
On behalf of the Virgo collaboration



Outline



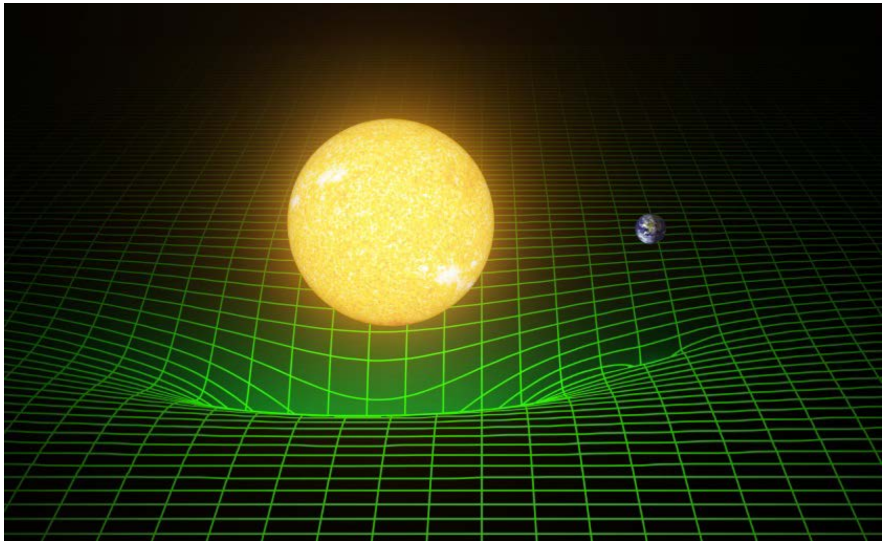
- Gravitational waves in a nutshell
  - Sources and properties
- Gravitational wave interferometric detectors
  - Principle and main characteristics
  - A worldwide network of detectors
- From Virgo to Advanced Virgo
  - Goals & upgrade
  - Status & plans
- The Advanced LIGO « Observation 1 » Run: September 2015 – January 2016
  - Performance
  - GW150914: the first direct detection of gravitational waves / black holes
- Outlook

Thanks to the many colleagues
from the LAL Virgo group, from Virgo and LIGO
from wich I borrowed ideas and material for this talk

Gravitational waves: sources and properties

General relativity in a nutshell

- “Spacetime tells matter how to move; matter tells spacetime how to curve”
  John Archibald Wheeler (1990)
  - A massive body warps the spacetime fabric
  - Objects (including light) move along paths determined by the spacetime geometry



- Einstein’s equations

$$G_{\mu\nu} = \frac{8\pi G}{c^4} T_{\mu\nu}$$

→ In words: Curvature = Matter

- Einstein tensor $G_{\mu\nu}$: manifold curvature
- Stress-energy tensor $T_{\mu\nu}$: density and flux of energy and momentum in spacetime
- Equality between two tensors
  → Covariant equations
- Need to match Newton’s theory for weak and slowly variable gravitational fields
  → Very small coupling constant: the spacetime is very rigid
- Non linear equations: gravitational field present in both sides

Schwartzschild Radius

- Newtonian escape velocity: $\mathbf{v_e = \sqrt{\frac{2GM}{r}}}$
- Schwartzschild radius R_S (1916): $\mathbf{R_S = \frac{2GM}{c^2} \approx 3km \left(\frac{M}{M_{Sun}}\right)}$
  - $R_S(M)$ such as $v_e = c$

→ Very small for « usual » celestial objects

    - Planets, stars

- Compacity $\mathbf{C = \frac{R_S}{radius} \leq 1}$

Object	**Earth**	**Sun**	**White dwarf**	**Neutron star**	**Black hole**
Compacity	**$1.4\ 10^{-9}$**	**$4.3\ 10^{-6}$**	**10^{-4}**	**0.3**	**1**

- Beware: compact and dense are two different things!
  - Black hole « density »

$$\boldsymbol{\rho = \frac{\text{"Mass"}}{\text{"Volume"}} \approx 1.8 \times 10^{16}\,\mathrm{g/cm^3} \left(\frac{M_{Sun}}{M}\right)^2}$$

Black holes

- Spacetime region in which gravitation is so strong that nothing, not even light, can escape from inside its horizon

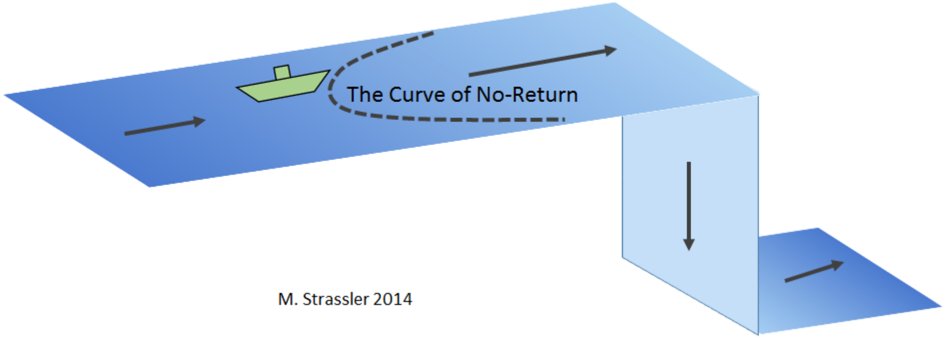
- Formed by the collapse of massive stars running out of fuel

- Can grow by accreting matter
  - Supermassive black holes are though to exist inside most galaxies
    → E.g. Sagittarius A* in the center of the Milky Way

- Characterized by three numbers (Kerr, 1963)
  - Mass
  - Spin
  - Electric charge

- Black hole horizon
  - Once crossed there’s no way back
  - Can only grow with time

A Person In a Boat that Crosses the Curve of No-Return Will Notice Nothing at the Time, But is Doomed To Go Over The Waterfall



M. Strassler 2014

Gravitational waves (GW)

- One of the first predictions of general relativity (1916)
  - Accelerated masses induce perturbations of the spacetime which propagate at the speed of light
  - Linearization of the Einstein equations ($g_{\mu\nu} = \eta_{\mu\nu} + h_{\mu\nu}$, $|h_{\mu\nu}| << 1$) leads to a propagation equation far from the sources

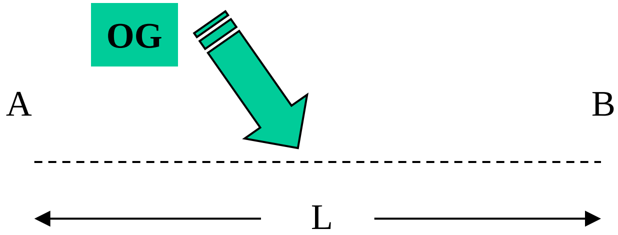
- Traceless and transverse (tensor) waves
  - 2 polarizations: « + » and « × »
    → See next slide for the interpretation of these names

- Quadrupolar radiation
  - Need to deviate from axisymmetry to emit GW
  - No dipolar radiation – contrary to electromagnetism

- GW amplitude h is dimensionless
  - Scales with the inverse of the distance from the source
  - GW detectors sensitive to amplitude ($h \propto 1/d$) and not intensity ($h^2 \propto 1/d^2$)
    → Important to define the Universe volume a given detector is sensitive to

Effect of gravitational waves on test masses

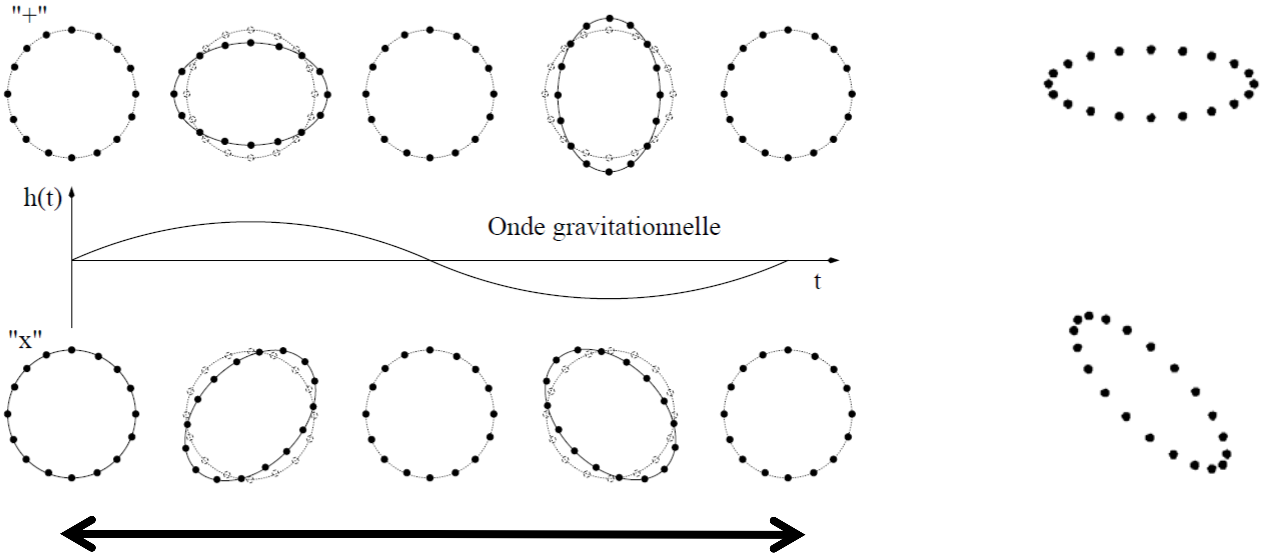
- GW: propagating perturbation of the spacetime metric
  - Acts on distance measurement between test masses (free falling)



$$\delta L_{max} = \frac{hL}{2}$$

Variation doubled for an interferometer with arms of equal length L:
$\delta L_{IFO} = hL$

- Effect of the two GW polarizations on a ring of free masses
  - « + » polarization
  - « × » polarization



One period

Do gravitational waves exist?

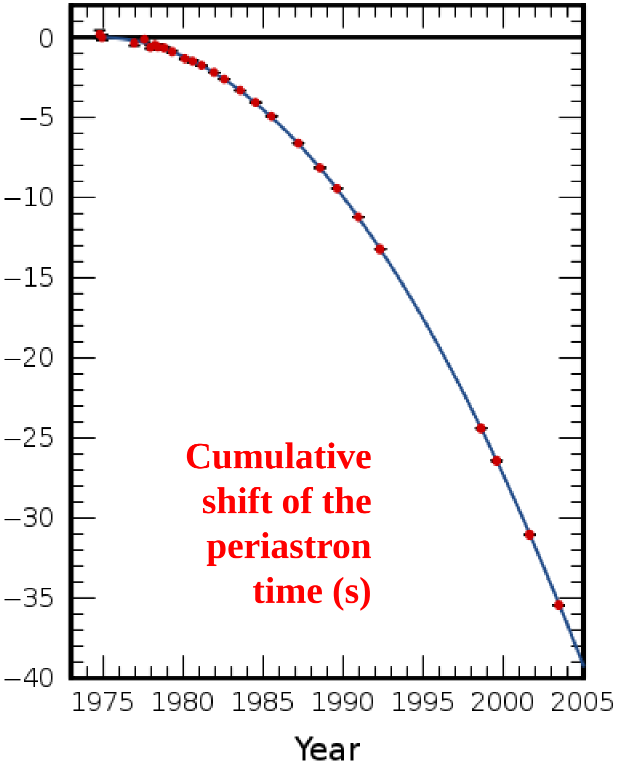
- Question (officially) solved since February 11 2016!
  - But was very relevant beforehand … and long-standing in the community

- Controversy for decades
  - Eddington, 1922: « *GW propagate at the speed of thought* »
  - 1950's: general relativity is mathematically consistent (Choquet-Buhat)

- Indirect evidence of the GW existence: long-term study of PSR B1913+16 – see next slide
  - Galactic (6.4 kpc away) binary system
  - Two neutron stars, one being a pulsar

- Discovered by Hulse and Taylor in 1974
  - Nobel prize 1993

- Laboratory for gravitation study
  - GW in particular
    → Taylor & Weisberg, Damour



Sources of gravitational waves

- Einstein quadrupole formula (1916)
  - Power radiated into gravitational waves

    Q: reduced quadrupole momenta

Very small: 10^{-53} W^{-1}

$$P = \frac{G}{5c^5} \left\langle \dddot{Q}_{\mu\nu} \, \dddot{Q}^{\mu\nu} \right\rangle$$

- Let's rewrite this equation introducing some typical parameters of the source
  - Mass M, dimension R, frequency $\omega/2\pi$ and asymmetry factor a
  - One gets $\dfrac{d^3Q}{dt^3} \sim \left(aMR^2\right)\omega^3$ and $P \sim \dfrac{G}{c^5} a^2 M^2 R^4 \omega^6$
- Using $\omega \sim v/R$ and introducing R_S, one gets:

$$\boxed{P \sim \frac{c^5}{G} a^2 \; C^2 \left(\frac{v}{c}\right)^6}$$

Huge: 10^{53} W

© Joe Weber, 1974

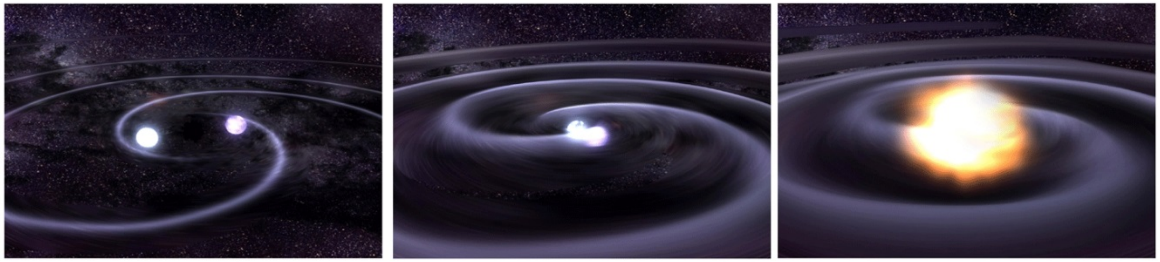
→ A good GW source must be

- Asymmetric
- As compact as possible
- Relativistic

- Although all accelerated masses emit GW, no terrestrial source can be detected

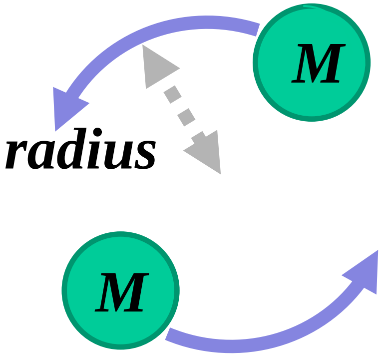
  → Need to look for astrophysical sources (typically: $h \sim 10^{-22} \div 10^{-21}$)

A diversity of sources



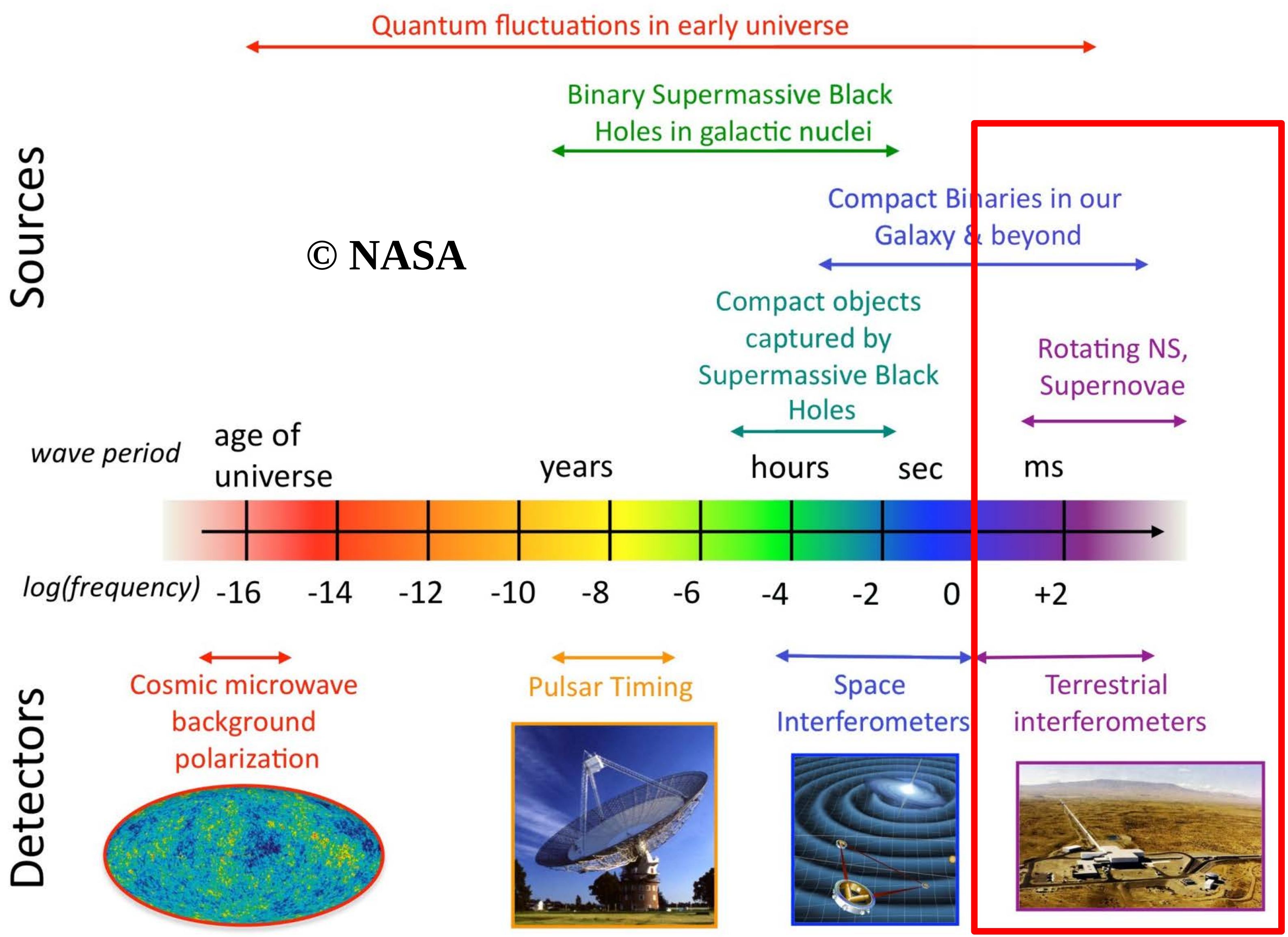
- Rough classification
  - Signal duration
  - Frequency range
  - Known/unknown waveform
  - Any counterpart (E.M., neutrinos, etc.) expected?
- Compact binary coalescence
  - Last stages of the evolution of a system like PSRB 1913+16
    → Compact stars get closer and closer while loosing energy through GW
  - Three phases: inspiral, merger and ringdown
    → Modeled via analytical computation and numerical simulations
  - Example: two masses M in circular orbit ($f_{GW} = 2\ f_{Orbital}$)

$$h \approx 10^{-21} \left(\frac{500\ \mathrm{Mpc}}{\mathrm{Distance}}\right) \left(\frac{\mathrm{Mass}}{30\ M_{Sun}}\right) \left(\frac{\mathrm{Orbital\ radius}}{100\ \mathrm{km}}\right)^2 \left(\frac{\mathrm{Frequency}}{100\ \mathrm{Hz}}\right)^2$$



- Transient sources (« bursts »)
  - Example: core collapses (supernovae)
- Permanent sources
  - Pulsars, Stochastic backgrounds

Gravitational wave spectrum



© NASA

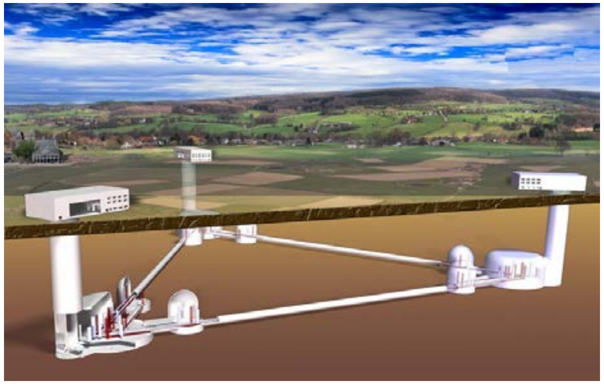
LIGO, Virgo, etc.

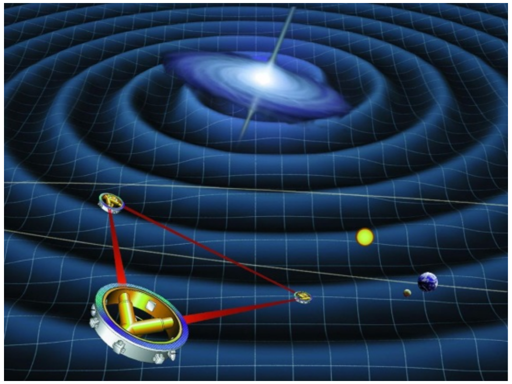
Gravitational wave detectors

EXPLORER resonant bar operated @ CERN 1991-2012

- On the ground
  - Resonant bars (Joe Weber's pioneering work)
    → Narrow band, limited sensitivity
  - Interferometric detectors
    → LIGO, Virgo and others
    → 2nd generation (« advanced ») detectors started operation
    Design studies have started for 3rd generation detectors (Einstein Telescope)
  - Pulsar Timing Array (http://www.ipta4gw.org)
    → GW would vary the time of arrival pulses emitted by millisecond pulsars

- In space
  - Future mission eLISA (https://www.elisascience.org, 2030's)
  - Technologies tested by the LISA pathfinder mission, sent to space last December





Gravitational wave interferometric detectors

1916-2016: a century of progress

- **1916: GW prediction (Einstein)**

1957 Chapel Hill Conference (Bondi, Feynman, Pirani, etc.)

Theoretical developments

- **1963: rotating BH solution (Kerr)**
- **1990's: CBC PN expansion (Blanchet, Damour, Deruelle, Iyer, Will, Wiseman, etc.)**
- **2000: BBH effective one-body approach (Buonanno, Damour)**
- **2006: BBH merger simulation (Baker, Lousto, Pretorius, etc.)**

Experiments

- **1960's: first Weber bars**
- **1970: first IFO prototype (Forward)**
- **1972: IFO design studies (Weiss)**
- **1974: PSRB 1913+16 (Hulse & Taylor)**
- **1980's: IFO prototypes (10m-long) (Caltech, Garching, Glasgow, Orsay)**
- **End of 1980's: Virgo and LIGO proposals**
- **1990's: LIGO and Virgo funded**
- **2005-2011: initial IFO « science » » runs**
- **2007: LIGO-Virgo Memorandum Of Understanding**
- **2012 : Advanced detectors funded**
- **2015: First Advanced LIGO science run**

Gravitational wave interferometric detectors

- Instructions to build a GW detector
  - Use free test masses
  - Locate them far apart
  - Measure their relative displacement
  - Make sure their motion is not perturbated by any external source

- Solution: a Michelson interferometer

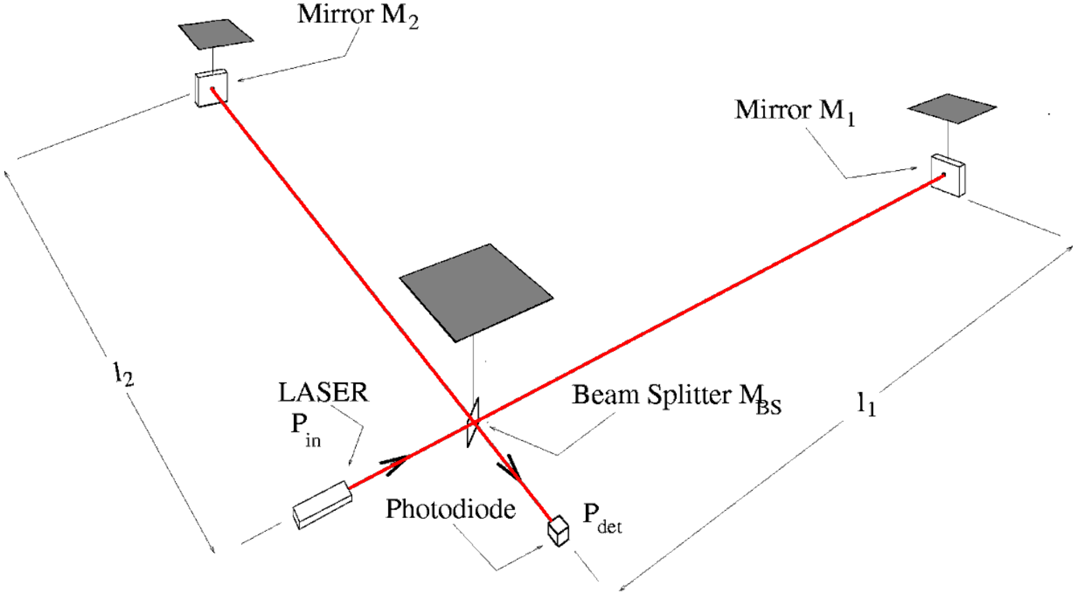
  → Suspended mirrors

  → Kilometer-long arms

  → Get rid of common mode noise

  → Design + active control

  + noise mitigation/monitoring



- Incident GW

  ⇒ Optical path changes

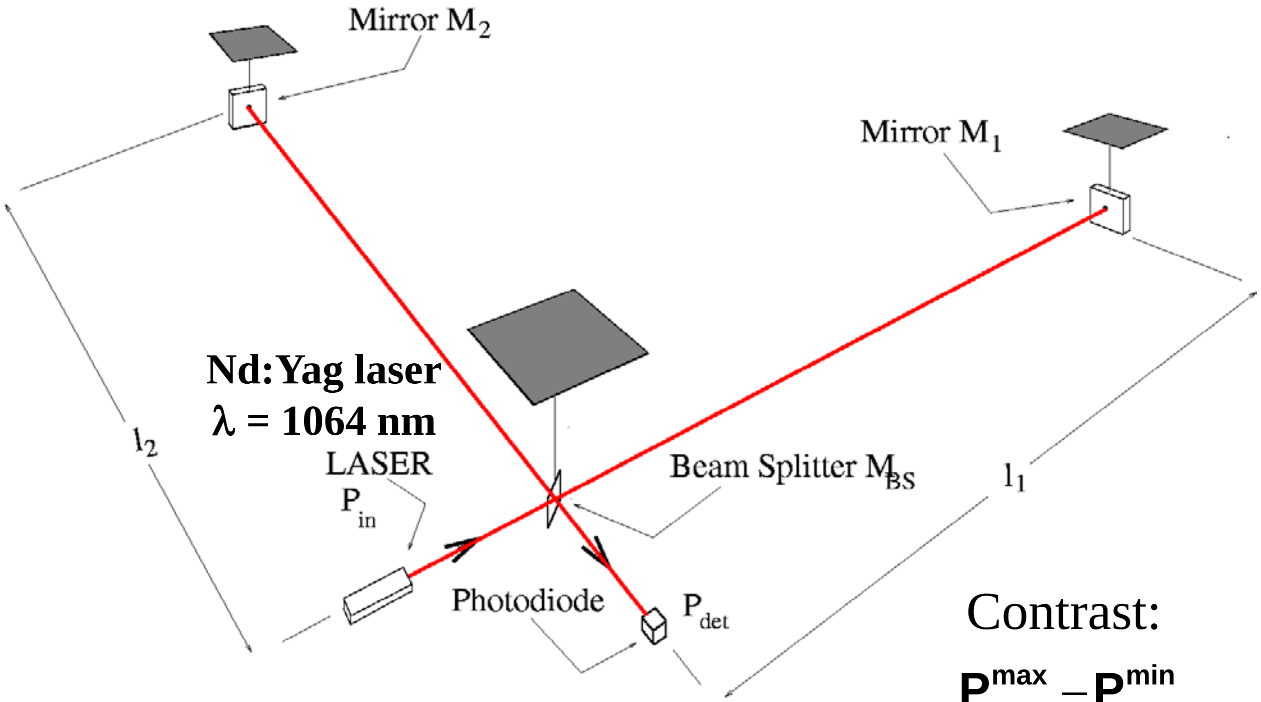
  ⇒ Output power variation

- Best sensitivity around the dark fringe

Suspended Michelson interferometer

- Mirrors act as test masses

- Incident GW
  $\rightarrow$ Modification of optical paths
  $\rightarrow$ Variation of detected light power



Contrast:

$$C = \frac{P^{max} - P^{min}}{P^{max} + P^{min}} \approx 1$$

- Output power

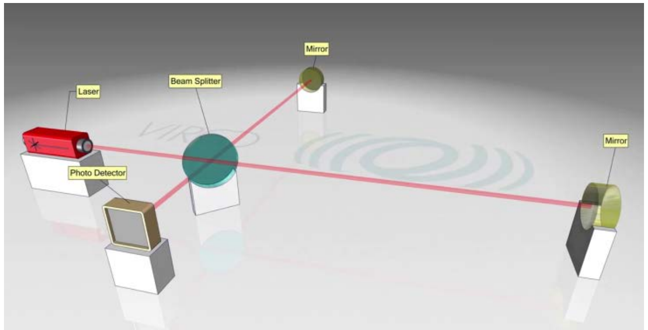
$$P_{det} = \frac{P_{in}}{2}\left[1 + C\cos(\Delta\phi)\right]$$

- Expanding the phase, one gets $\Delta\phi = \overset{\equiv \Delta\phi_{OP}}{\boxed{\frac{2\pi(l_2 - l_1)}{\lambda}}} + \overset{\equiv \delta\phi_{GW}}{\boxed{\frac{2\pi(l_1 + l_2)h(t)}{\lambda}}}$

- and finally $P_{det} \approx \frac{P_{in}}{2}\left[1 + C\cos(\Delta\phi_{OP}) - C\sin(\Delta\phi_{OP}) \times \boxed{\delta\phi_{GW}(t)}\right]$ Output power variation $\propto h(t)$

- Working point set ~10^{-11} m away from the dark fringe

Interferometer sensitivity



Suspensions not drawn

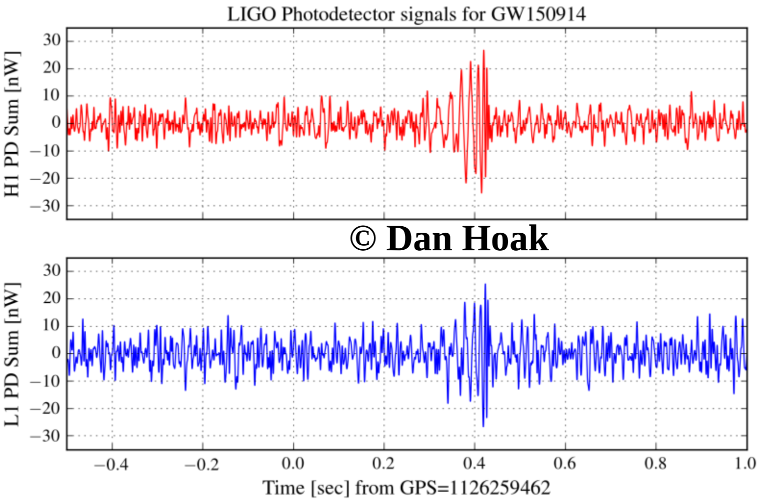
- Output power: $\delta \mathbf{P}_{\mathbf{det}} \propto \mathbf{P}_{\mathbf{in}}\, \mathbf{L}\, \mathbf{h}$
- Shot noise
  - A fundamental quantum noise
  - Fluctuation of the number of photons detected during a duration Δt

$$\delta \mathbf{P}_{\mathbf{shot\ noise}} \propto \sqrt{\frac{\mathbf{P}_{\mathbf{in}}}{\Delta \mathbf{t}}}$$

$$\delta \mathbf{P}_{\mathbf{det}} = \delta \mathbf{P}_{\mathbf{shot\ noise}}$$

- Minimum detectable GW amplitude such that

$$\rightarrow \boxed{\mathbf{h}_{\mathbf{min}} \propto \frac{1}{\sqrt{\mathbf{P}_{\mathbf{in}}}\; \mathbf{L} \sqrt{\Delta \mathbf{t}}}}$$



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- Improving the sensitivity
  - Increase incident power on the beamsplitter
  - Increase length of the interferometer arms

- Reaching h_{min}~10^{-22} or below requires
  - Kilowatts of laser power and
  - Arms about a hundred kilometer long

} Virgo/LIGO design

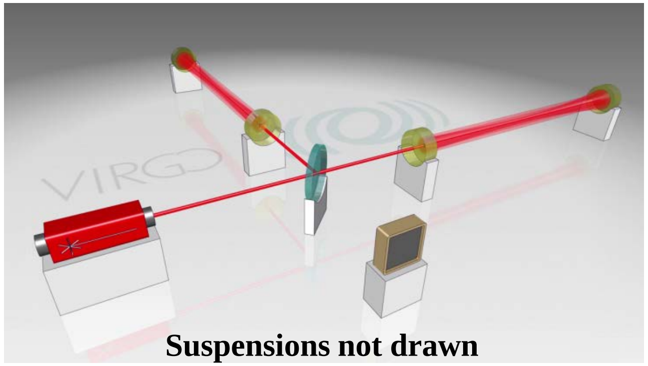
Bandpass and notch filtering
25 nW offset subtracted
500 W incident on the beamsplitter

Improving the interferometer sensitivity

- Reminder: Interferometer (IFO) sensitivity $\propto \dfrac{1}{(\text{Arm length}) \times \sqrt{\text{Light power}}}$

→ Use high power laser, power- and frequency-stabilized

- Tens to hundreds of watts



Suspensions not drawn

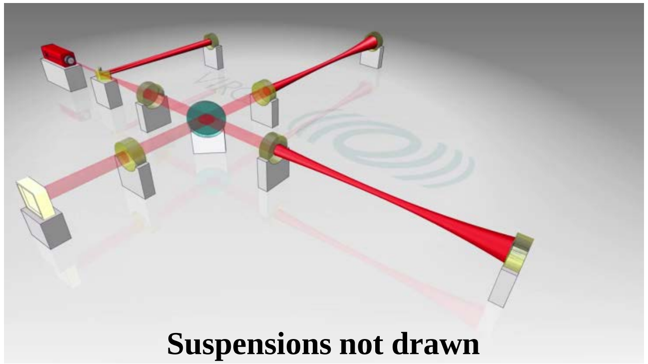
→ Kilometric arms (Virgo: 3km; LIGO: 4km)

→ Add Fabry-Perot cavities in the kilometric arms

- Light path length increased: $L \rightarrow L \times G_{FP}$
  G_{FP}~300 for Advanced Virgo
- Low-pass filter on the IFO frequency response: processes faster than the light storage time are filtered

→ Add recycling mirror between the input laser and the beamsplitter

- IFO set to the dark fringe + highly reflecting mirrors } All power reflected back to the laser!
  $P_{in} \rightarrow P_{in} \times G_{rec}$, G_{rec}~40 for Advanced Virgo

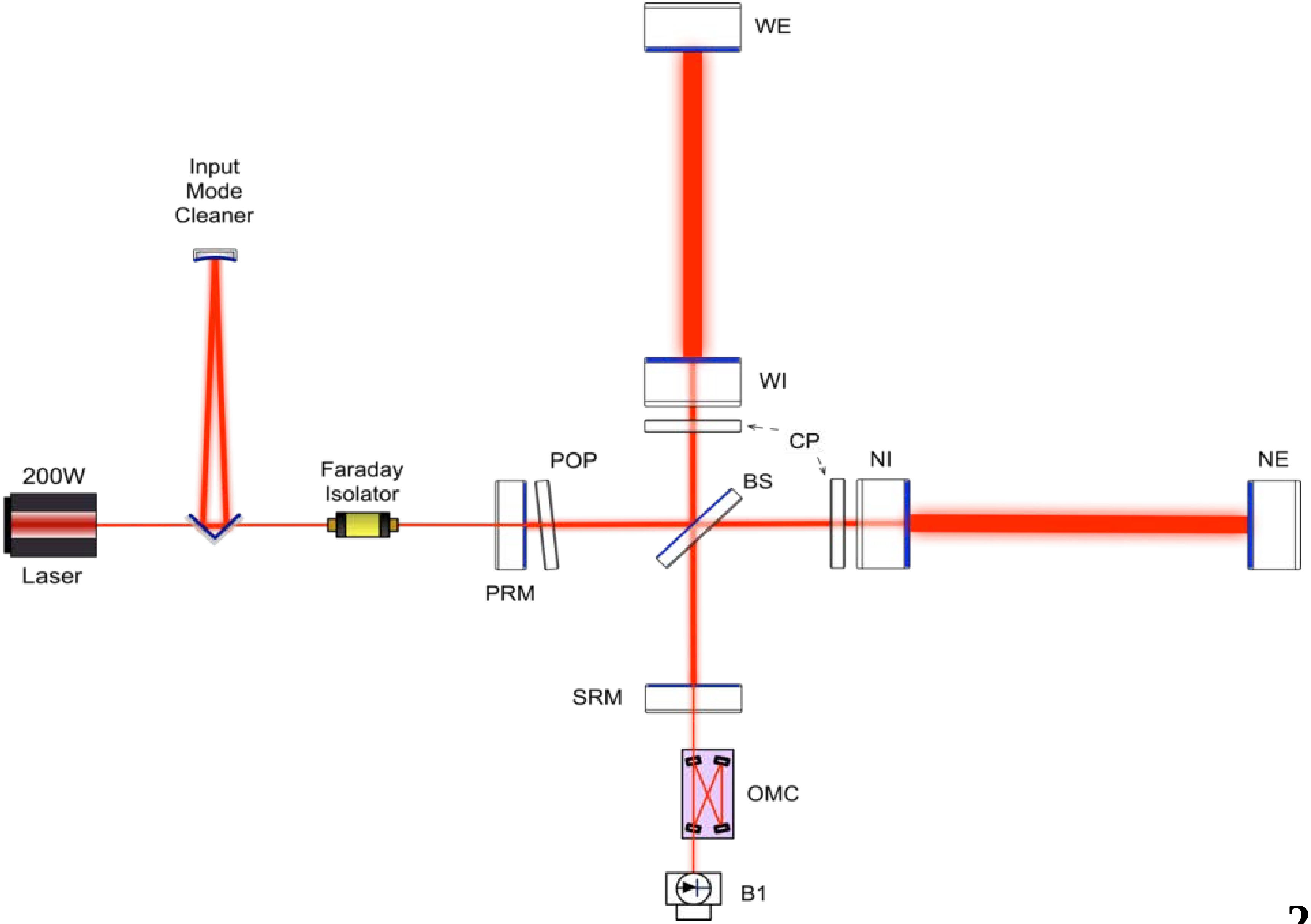


Suspensions not drawn

→ Minimize transmission and losses for all mirrors

- Set the gains of the interferometer cavities

The Advanced Virgo detector scheme



Noise & sensitivity

- Noise: any kind of disturbance which pollutes the dark fringe output signal
- Detecting a GW of frequency f ↔ amplitude h « larger » than noise at that frequency
- Interferometers are wide-band detectors
  - GW can span a wide frequency range
  - Frequency evolution with time is a key feature of some GW signals

    → Compact binary coalescences for instance
- Numerous sources of noise
  - Fundamental

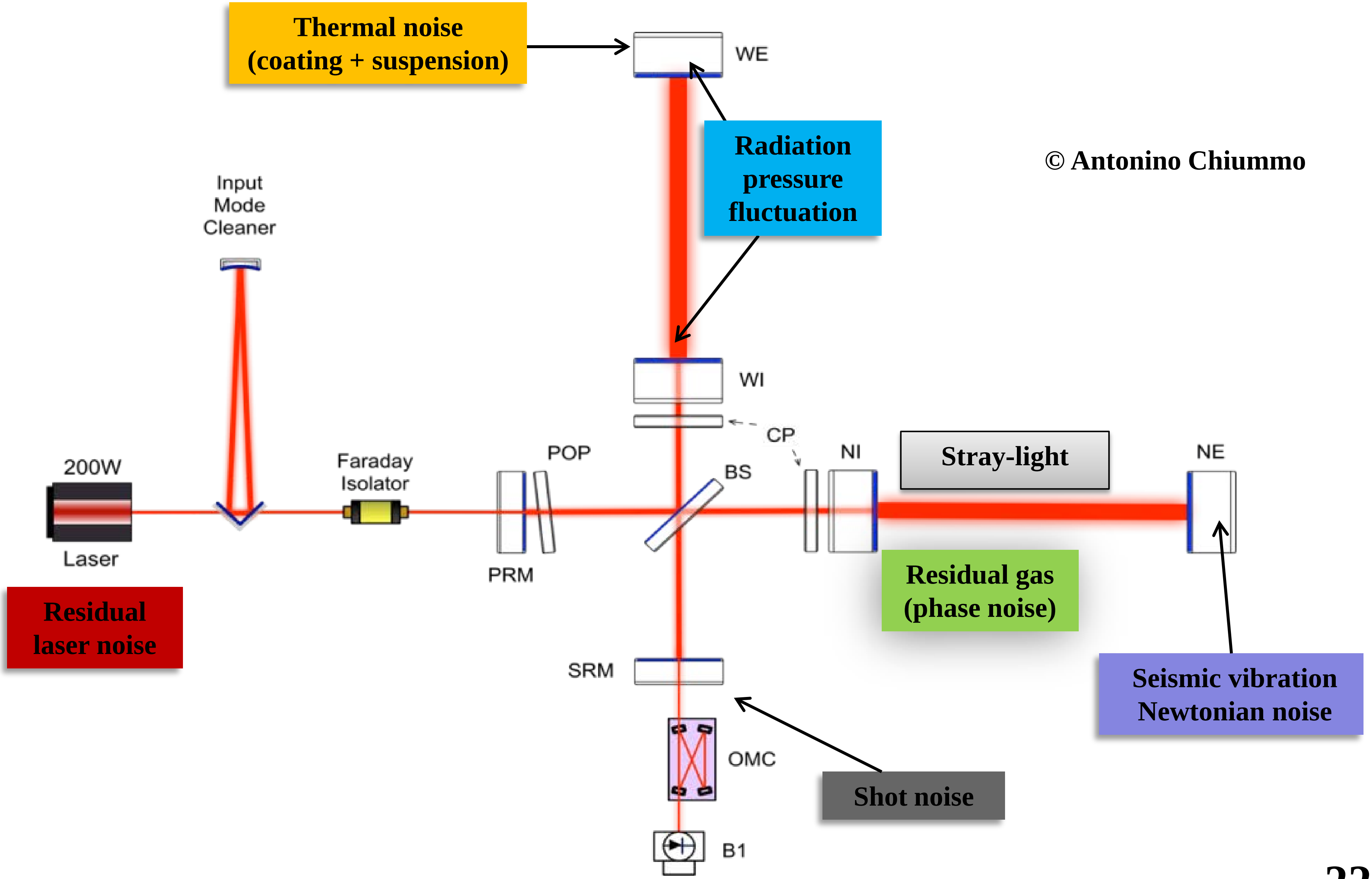
    → Cannot be avoided; optimize design to minimize these contributions
  - Instrumental

    → For each noise, identify the source; then fix or mitigate

    → Then move to the next dominant noise; iterate…
  - Environmental

    → Isolate the instrument as much as possible; monitor external noises
- IFO sensitivity characterized by its power spectrum density (PSD, unit: 1/√Hz)
  - Noise RMS in the frequency band $[f_{min};f_{max}] = \sqrt{\int_{f_{min}}^{f_{fmax}} \mathrm{PSD}^2(f)\,df}$

Main interferometer noises



Thermal noise (coating + suspension)

Radiation pressure fluctuation

© Antonino Chiummo

Stray-light

Residual gas (phase noise)

Residual laser noise

Seismic vibration Newtonian noise

Shot noise

Interferometer control

- A complex working point
  - Resonant Fabry-Perot and recycling cavities + IFO on the dark fringe
  - Arm length difference controled with an accuracy better than 10^{-15} m
  - The better the optical configuration, the narrower the working point

- « Locking » the IFO is a non-trivial engineering problem
  - Use several error signals to apply corrections on mirror positions and angles
    - → Pound-Drever-Hall signals (phase modulation)
    - → Auxiliary green lasers (for 2nd generation IFOs)
  - Feedback loops from few Hz to few kHz
  - Cope with filter bandwith and actuator range

- Multi-step lock acquisition procedure

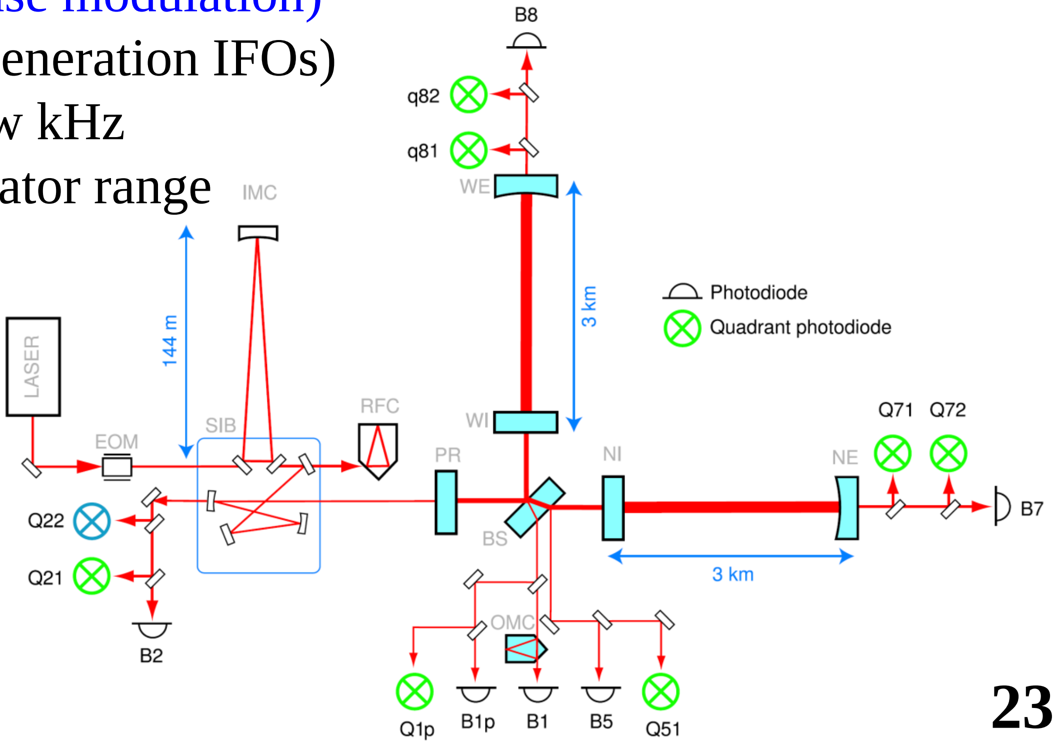
Free mirrors

↓

Local control

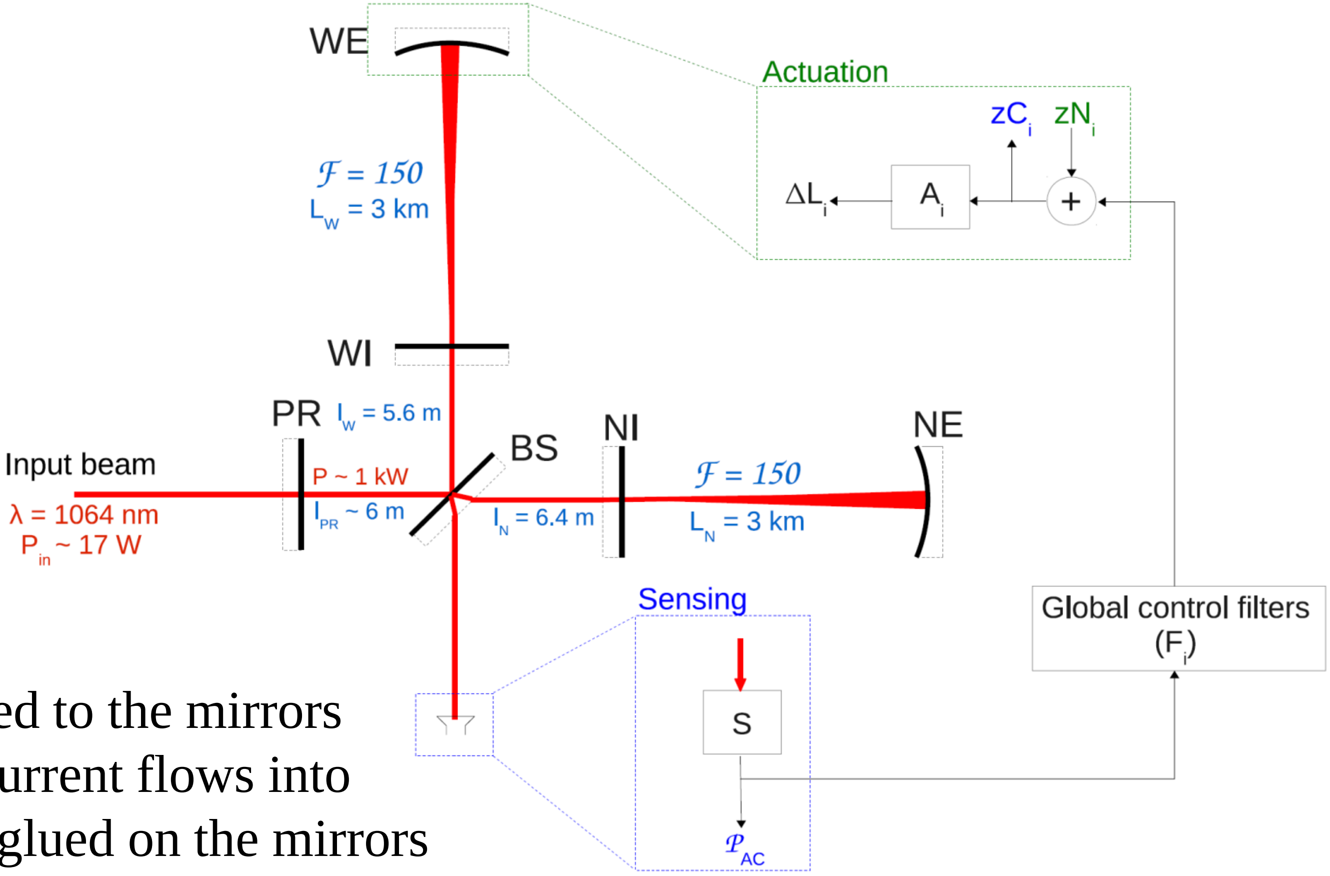
↓

Global control



Control chain

- Example of the dark fringe error signal
- Sensing
  - Photodiode readout
- Filtering
  - Algorithms use error signals to compute globally corrections sent to the mirrors
- Actuation
  - Corrections are applied to the mirrors by the suspensions: current flows into coils facing magnets glued on the mirrors



- Dedicated measurements to compute the sensing and actuation transfer functions

Reconstruction of the «GW channel»

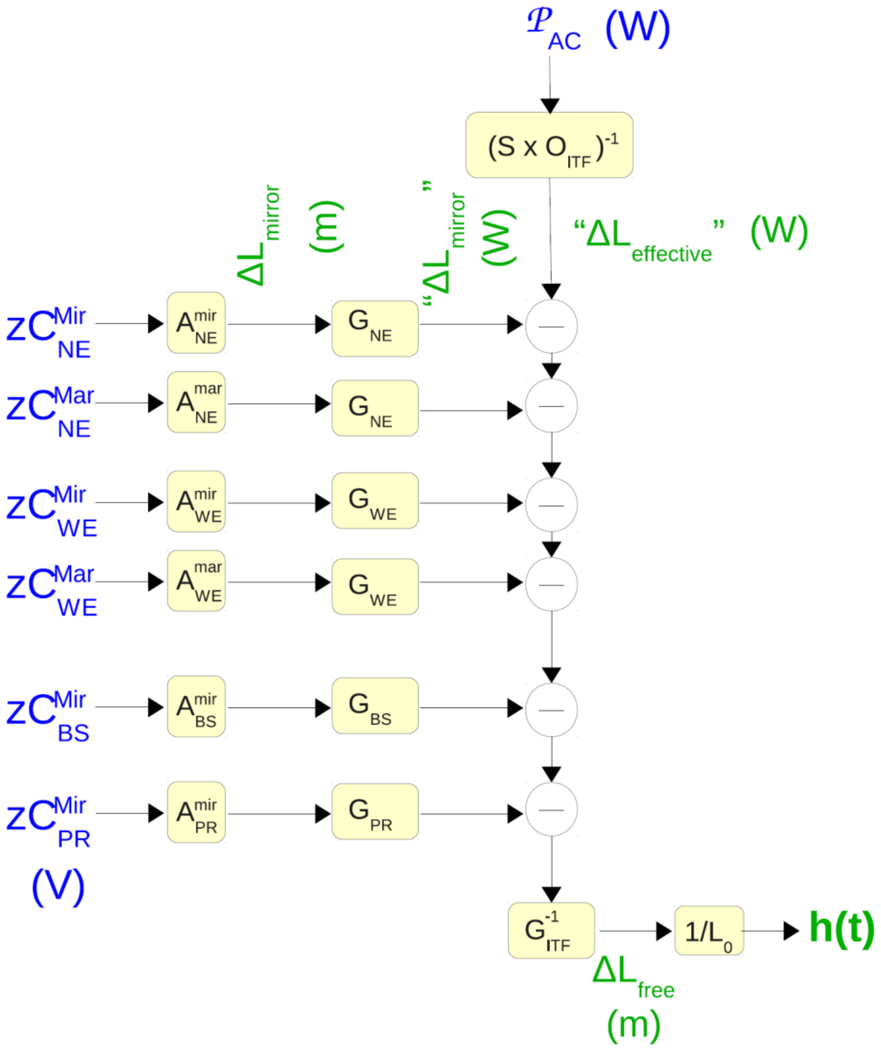
- Control loops act up to a few hundred Hz, both on noise and on a possible GW signal
  - Need to subtract their contributions to get h(t) = noise(t) [+ possibly GW(t)]

- Cavity optical transfer functions (W/m) directly measured by acting on mirrors during dedicated runs
  - Laser wavelength used as benchmark: → Frequency known at the Hz level

$$\Delta\phi = \frac{4\pi\Delta L}{\lambda}$$

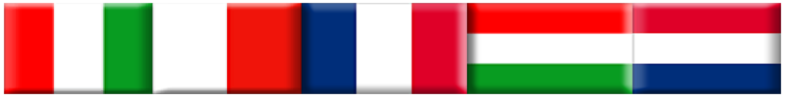
- Various gains monitored using calibration lines injected on each mirror suspension

- Finally, divide by the arm length to get h(t)



The Virgo collaboration

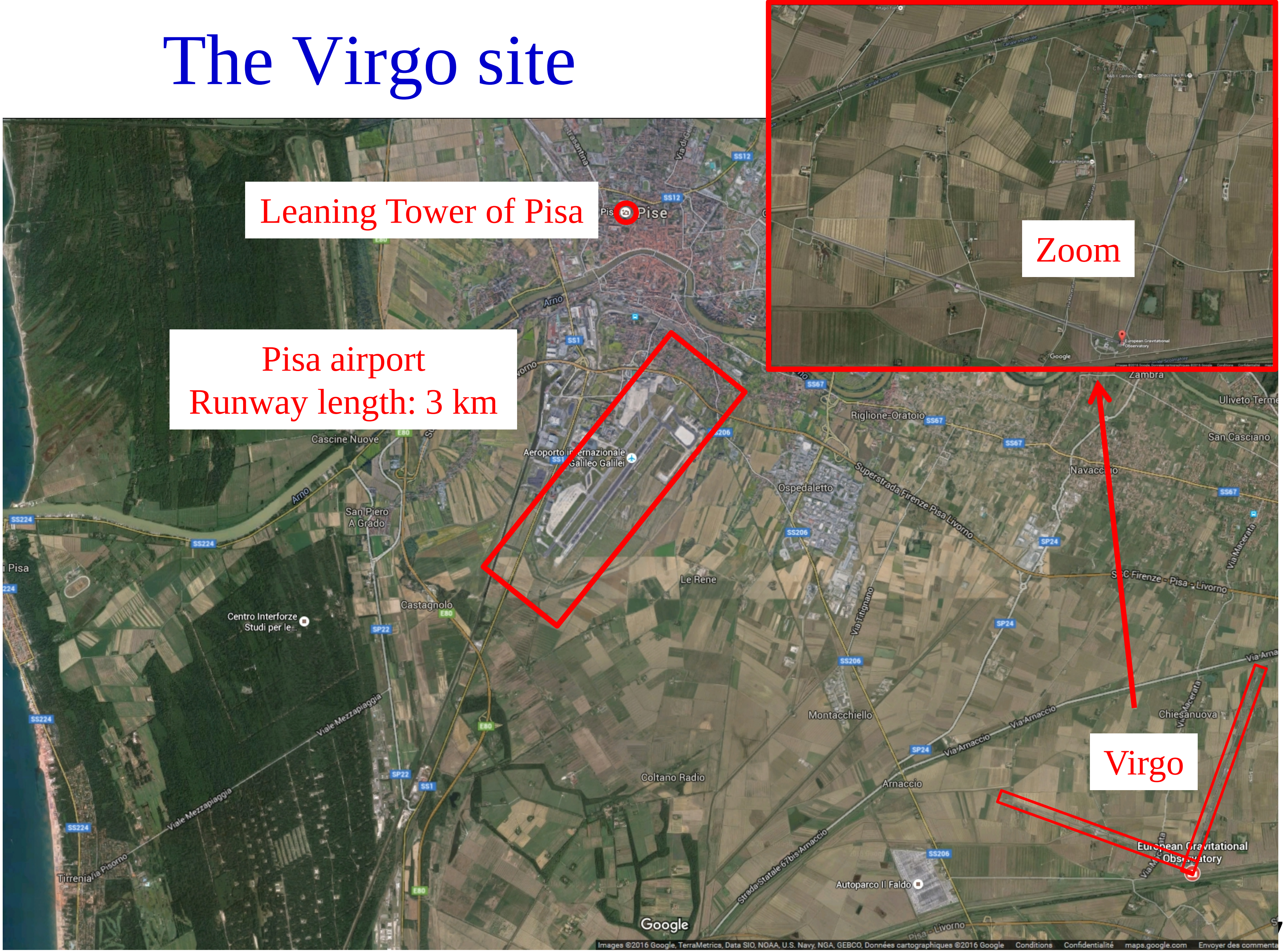
- 5 European countries



- 20 laboratories
- About 250 members (LIGO: 750)
- Virgo was built by 11 CNRS (France) and INFN (Italy) laboratories
  - Budget: ~150 M€
  - Groups from the Netherlands, Poland and Hungary joined later the project
- Advanced Virgo funding: ~20 M€
  - Plus in-kind contribution from NIKHEF
- The EGO (European Gravitational Observatory) consortium is managing the Virgo site in Cascina. It provides the infrastructures and ressources to ensure the detector construction and operation

APC Paris
ARTEMIS Nice
EGO Cascina
INFN Firenze-Urbino
INFN Genova
INFN Napoli
INFN Perugia
INFN Pisa
INFN Roma La Sapienza
INFN Roma Tor Vergata
INFN Padova
INFN TIFPA
LAL Orsay – ESPCI Paris
LAPP Annecy
LKB Paris
LMA Lyon
NIKHEF Amsterdam
POLGRAW (Poland)
RADBOUD Uni. Nijmegen
RMKI Budapest

The Virgo site



Leaning Tower of Pisa

Pisa airport
Runway length: 3 km

Zoom

Virgo

Network of gravitational wave interferometric detectors

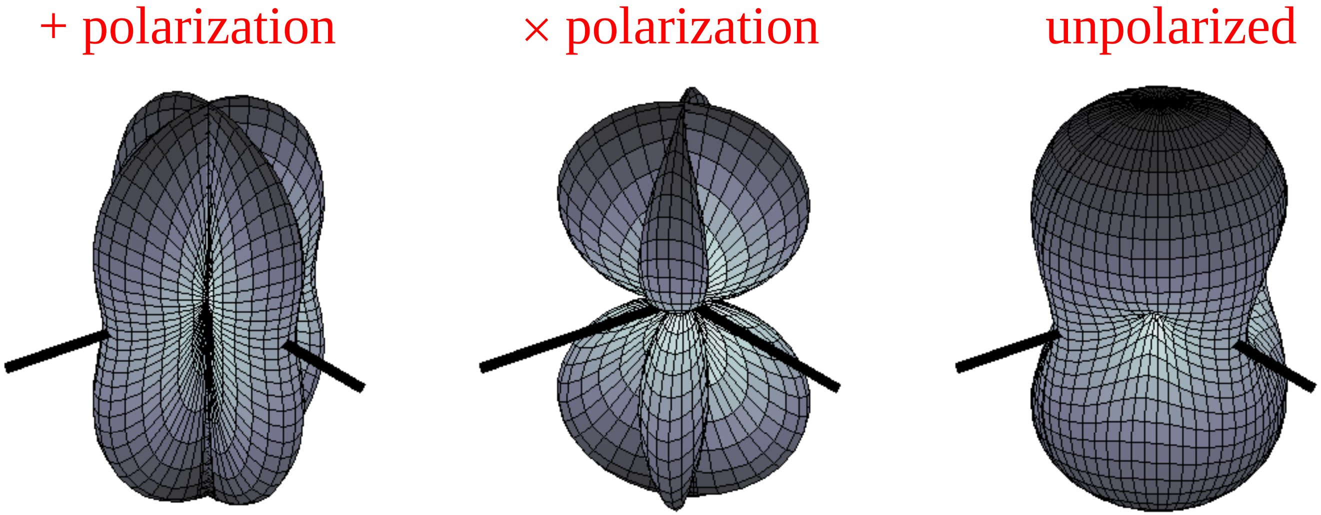
Interferometer angular response

- An interferometer is not directional: it probes most of the sky at any time
  - More a microphone than a telescope!

- The GW signal is a linear combination of its two polarisations

$$h(t) = F_+(t) \times h_+(t) + F_\times(t) \times h_\times(t)$$

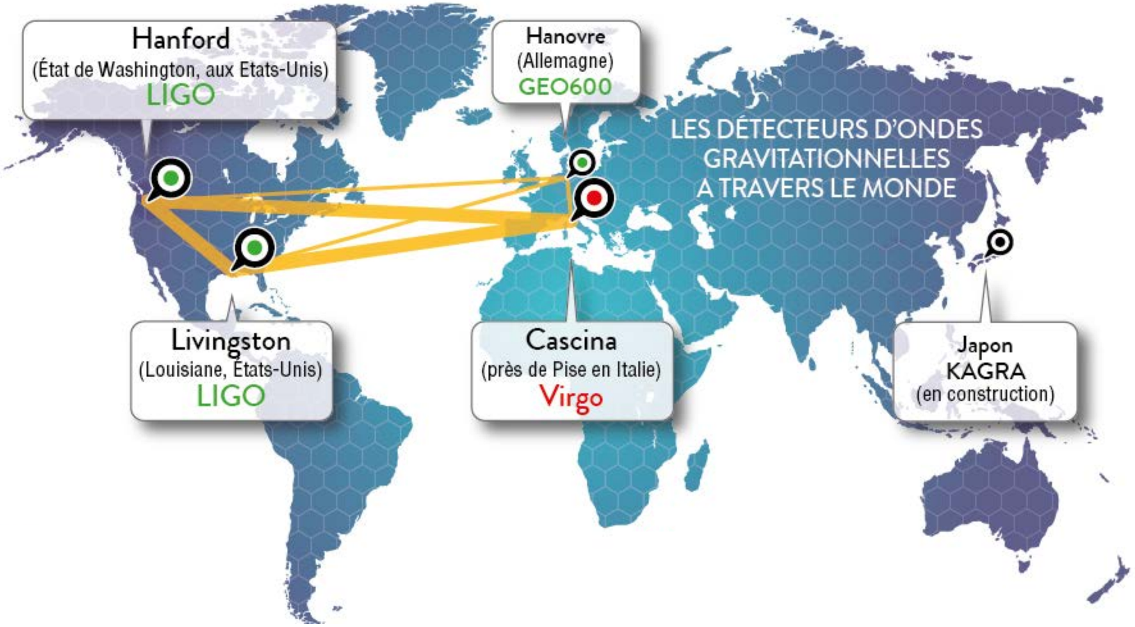
  - F_+ and $F_\times$ are antenna pattern functions which depend on the source direction in the sky w.r.t. the interferometer plane
    → Maximal when perpendicular to this plane
    → Blind spots along the arm bisector (and at 90 degres from it)



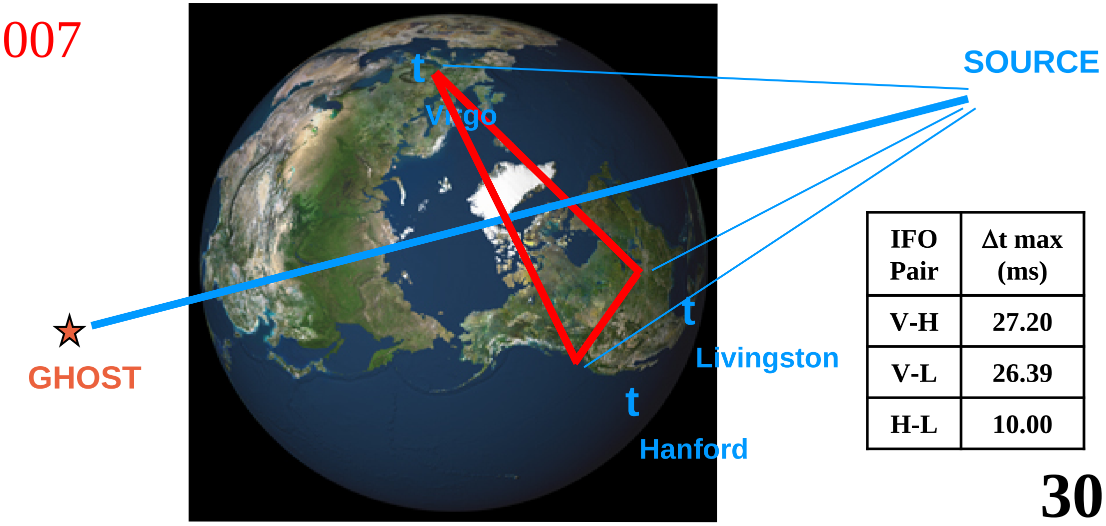
A network of interferometric detectors

- A single interferometer is not enough to detect GW
    - Difficult to separate a signal from noise confidently
    - There have been unconfirmed claims of GW detection

→ Need to use a network of interferometers



- Agreements (MOUs) between the different projects – Virgo/LIGO: 2007
    - Share data, common analysis, publish together
- IFO: non-directional detectors; non-uniform response in the sky
- Threefold detection: reconstruct source location in the sky



IFO Pair	Δt max (ms)
V-H	27.20
V-L	26.39
H-L	10.00

A network of interferometric detectors

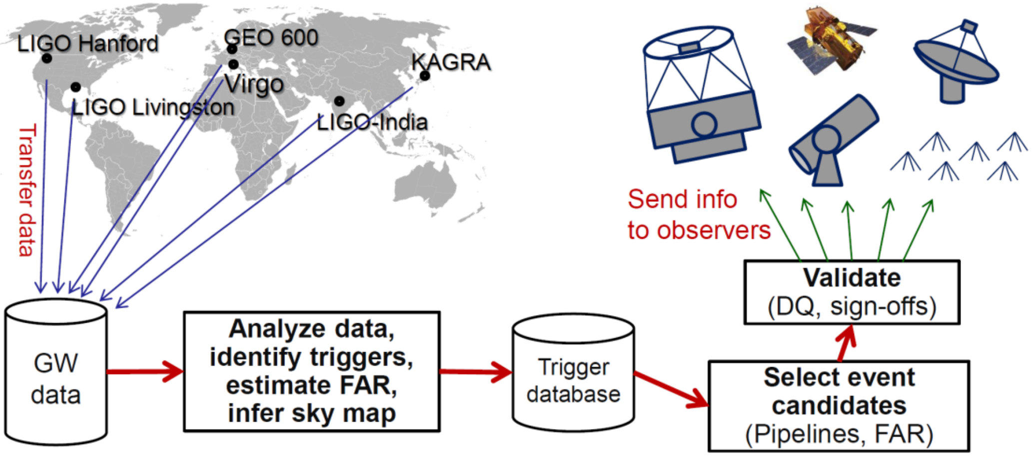

LIGO Hanford
Washington State, USA


LIGO Livingston
Louisiana, USA


Virgo Cascina (near Pisa), Italy

Exploiting multi-messenger information

- Transient GW events are energetic
  - Only (a small) part of the released energy is converted into GW
    → Other types of radiation released: electromagnetic waves and neutrinos
- Astrophysical alerts ⇒ tailored GW searches
  - Time and source location known ; possibly the waveform
    → Examples: gamma-ray burst, type-II supernova
- GW detectors are also releasing alerts to a worldwide network of telescopes
  - Agreements signed with ~75 groups – 150 instruments, 10 space observatories



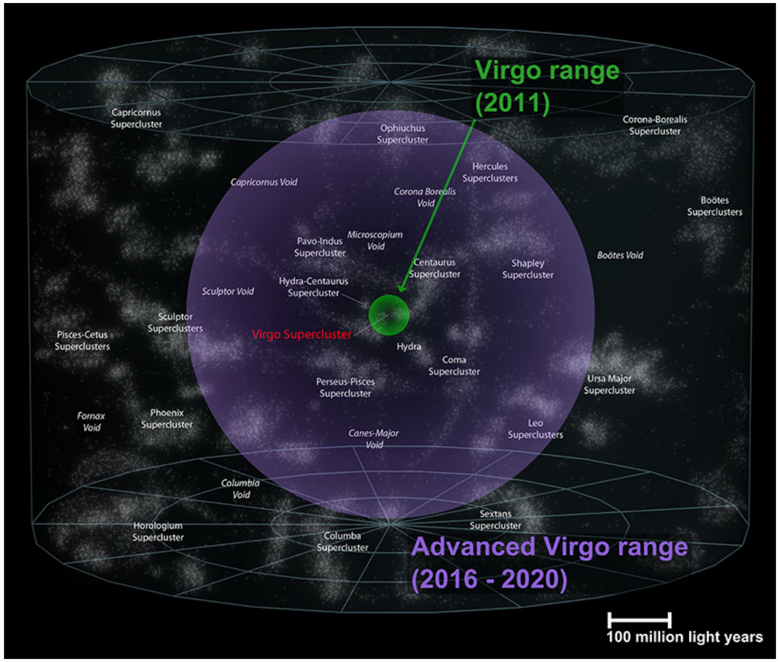
- Low latency h-reconstruction and data transfer between sites
  - Online GW searches for burst and compact binary coalescences

From Virgo to Advanced Virgo

From initial to advanced detectors

- Goal: to improve the sensitivity by one order of magnitude
  - Volume of observable Universe multiplied by a factor 1,000
  - Rate should scale accordingly
    → Assuming uniform distribution of sources (true at large scale)

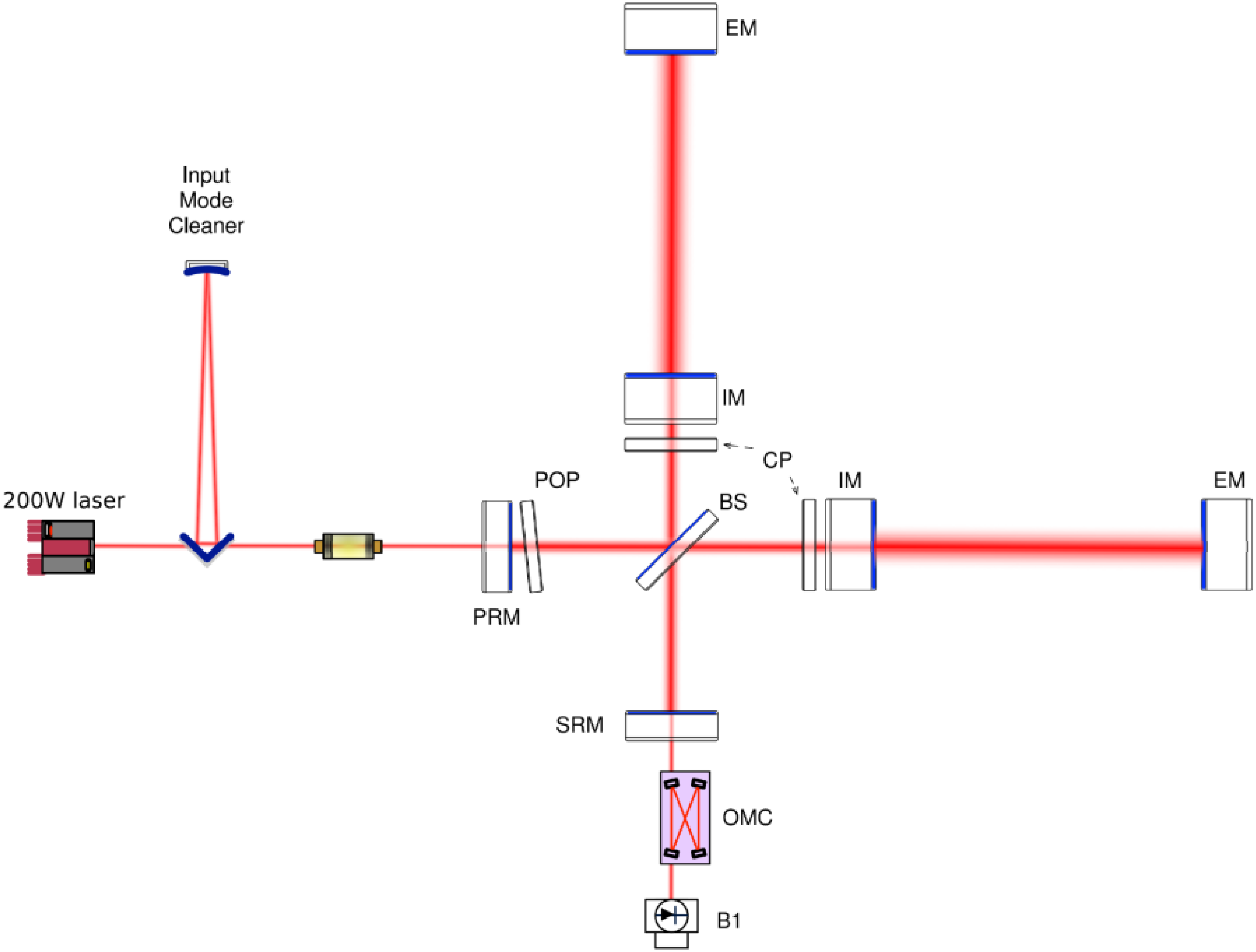
- A wide range of improvements
  - Increase the input laser power
  - Mirrors twice heavier
  - Increase the beamspot size on the end mirrors
  - Fused silica bonding to suspend the mirrors
  - Improve vacuum in the km-long pipes
  - Cryotraps at the Fabry-Perot ends
  - Instrumentation & optical benches under vacuum



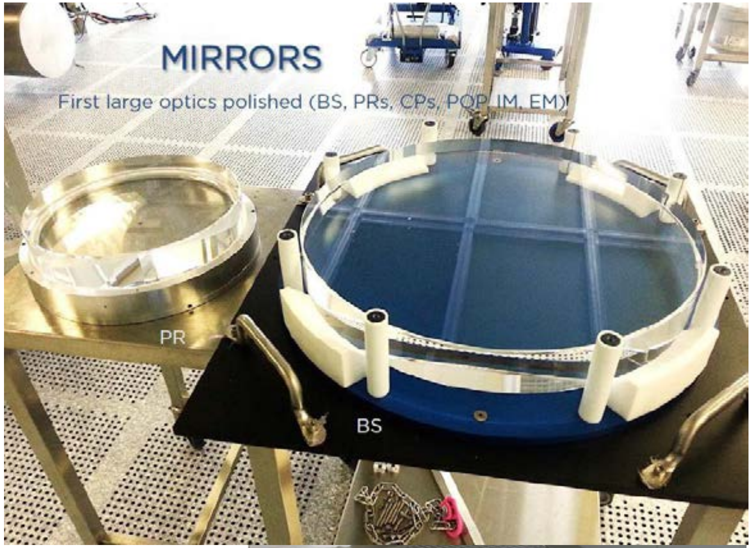
- Advanced LIGO (aLIGO) funded a year or so before Advanced Virgo (AdV)
  - Financial crisis in 2008-2010…

  → aLIGO ready for its first « observation run » in September 2015
  - AdV upgrade still in progress

The Advanced Virgo design



Mirrors



- SiO$_2$ substrates produced by Heraeus
- Coating in monoatomic layers performed at LMA (CNRS, Lyon)

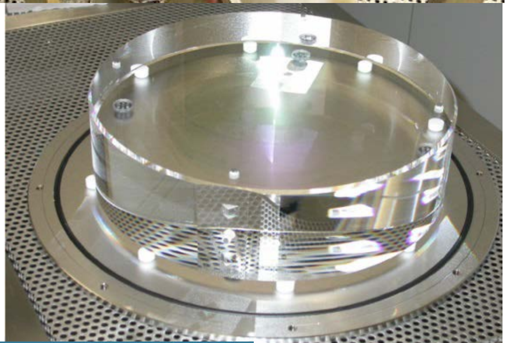
- Weight: few tens of kg, for a 35 cm diameter

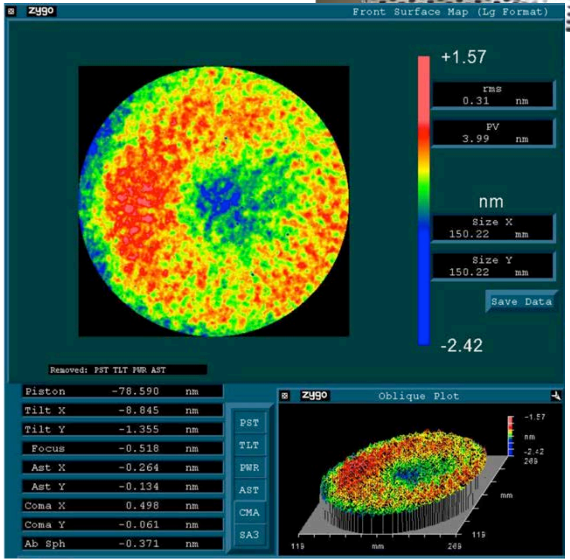
- Reflectivity set with an accuracy better than 0.1%
- Few ppm losses @ 1064 nm (nominal laser wavelength)
- Flatness below the nm over a 150 mm diameter
- Radius of curvature around 1500 m
  (half the long cavity length), accurate within a few meters
- Production completed on schedule

- Mirror measurements better than requirements
  - Less aberrations and scattered light

- Measured mirror maps included in Virgo simulations to predict the IFO behavior

- SiO$_2$ « ears » attached to the mirrors using an innovative silicate bonding technique





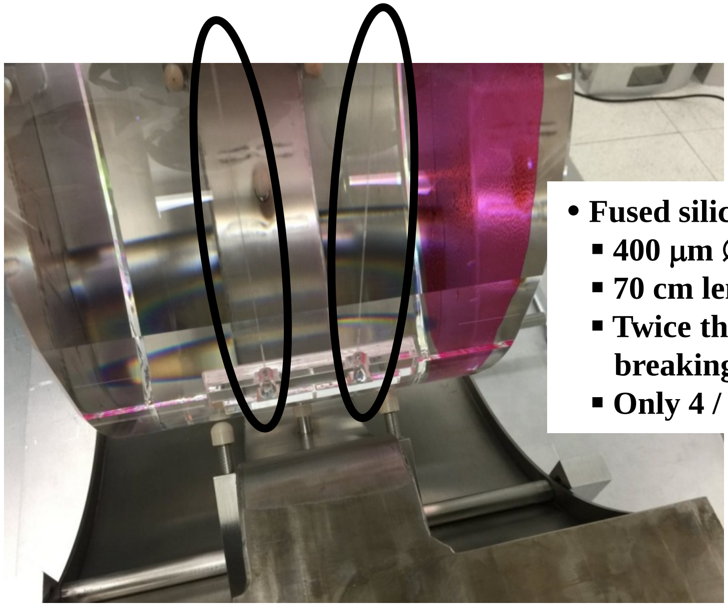
```
rms
0.31     nm
PV
3.99     nm
Size X
150.22   mm
Size Y
150.22   mm

Removed: PST TLT PWR AST

Piston    -78.590   nm
Tilt X     -8.845   nm
Tilt Y     -1.355   nm
 Focus     -0.518   nm
 Ast X     -0.264   nm
 Ast Y     -0.134   nm
Coma X      0.498   nm
Coma Y     -0.061   nm
Ab Sph     -0.371   nm
```

Low and medium frequency range improvements

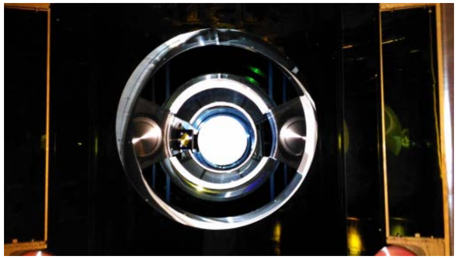
- Suspension and mirror thermal noises
    - Doubling the mirror weight (42 kg)
      → Noise scales like 1/√mass
    - Mirrors suspended with fused silica fibers
      → Smaller losses
    - Enlarging the beam size on the mirrors
      → Moving the beam waist close
      to the center of the long cavities
      → Larger vacuum links & beamsplitter
    - New low-dissipation mirror coatings



- **Fused silica fibers**
    - **400 μm ∅**
    - **70 cm length**
    - **Twice the steel breaking strength**
    - **Only 4 / mirror**

- Lowering the residual gas noise
    - Cryotraps at 77 K in between
      the towers and the 3 km-long tubes

- Limiting environmental noise
    - Photodiodes under vacuum on suspended benches
    - New baffles to fight stray light



High frequency range improvements

- Higher laser power
  - 125 W in the final configuration

  → New laser system

- Higher finesse in the Fabry-Perot cavities
  - Gain ~ 300: up to 700 kW stored

  → Very high-quality optics

  → Improved Thermal Compensation System (TCS)

- Signal recycling mirror to be added later in front of the dark port
  - Improve and shape the sensitivity curve in a given frequency band (tuning for specific sources)

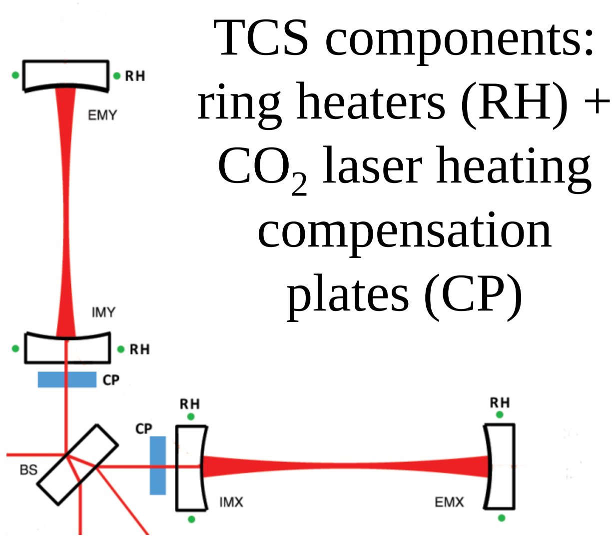
    Mirror reflectivity ↔ Bandwidth

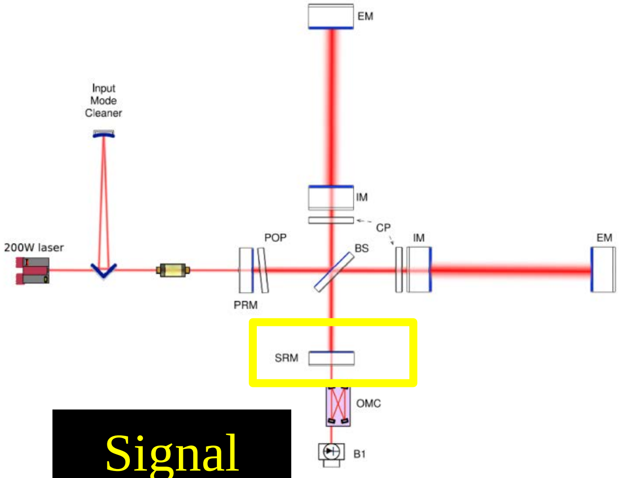
    Microscopic position ↔ Resonance frequency
  - Additional cavity to control

- DC detection at the dark port

  → New suspended optical benches

TCS components:
ring heaters (RH) +
CO_2 laser heating
compensation
plates (CP)

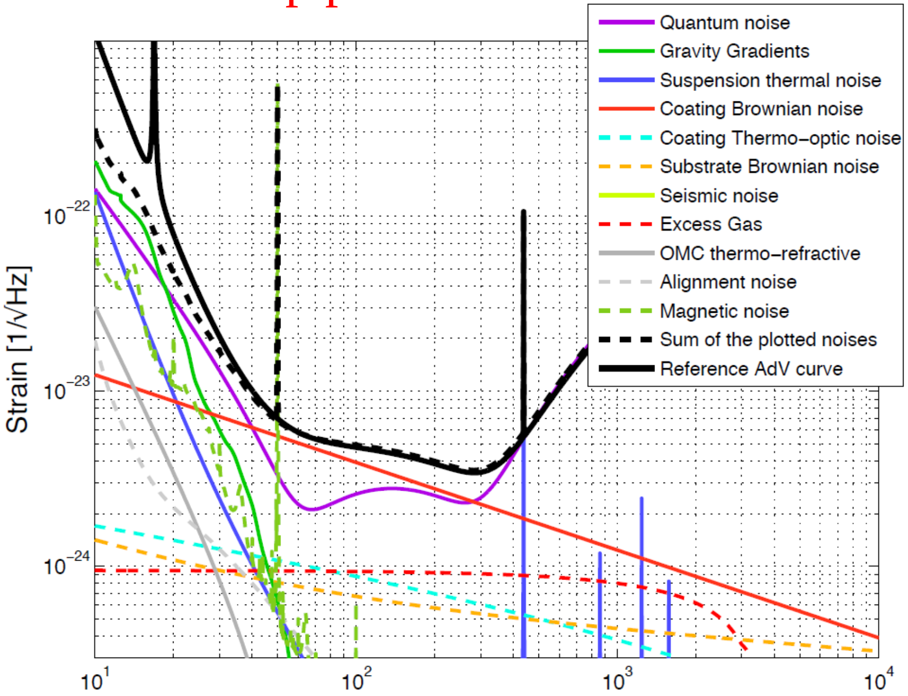


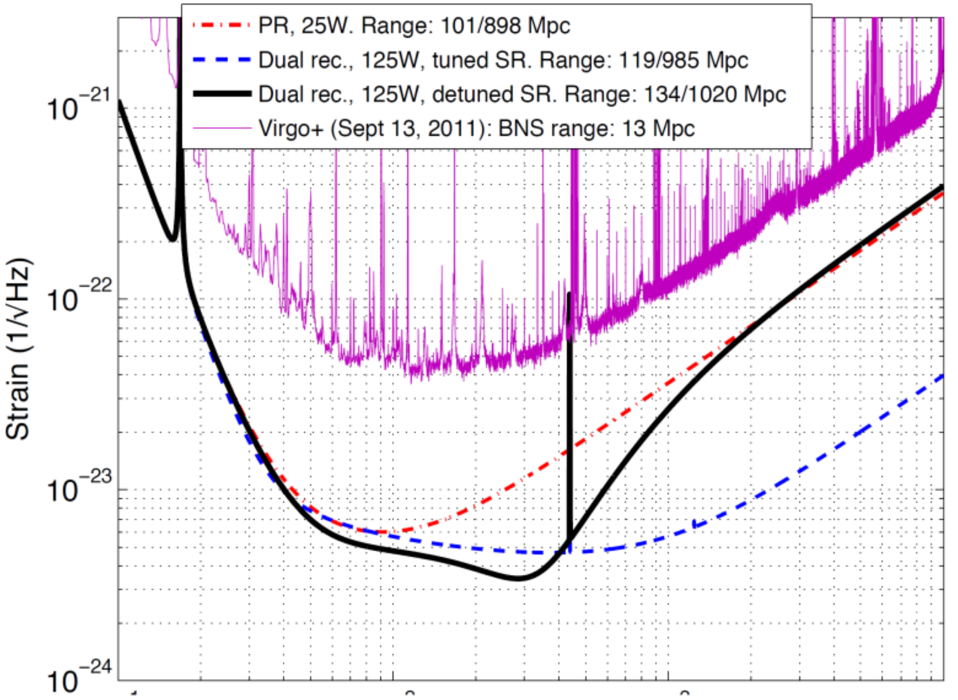


Signal
recycling
mirror

Sensitivity improvement

- A multi-step process





- Quantum noise dominant at low (radiation pressure) & high (shot noise) frequencies
  → R&D ongoing on frequency-dependent light squeezing
- Coating thermal noise dominant in between

- Low frequency sensitivity ultimately limited by Newtonian noise
  - Stochastic gravitational field induced by surface seismic waves
    → Either active cancellation or go underground

Advanced Virgo status

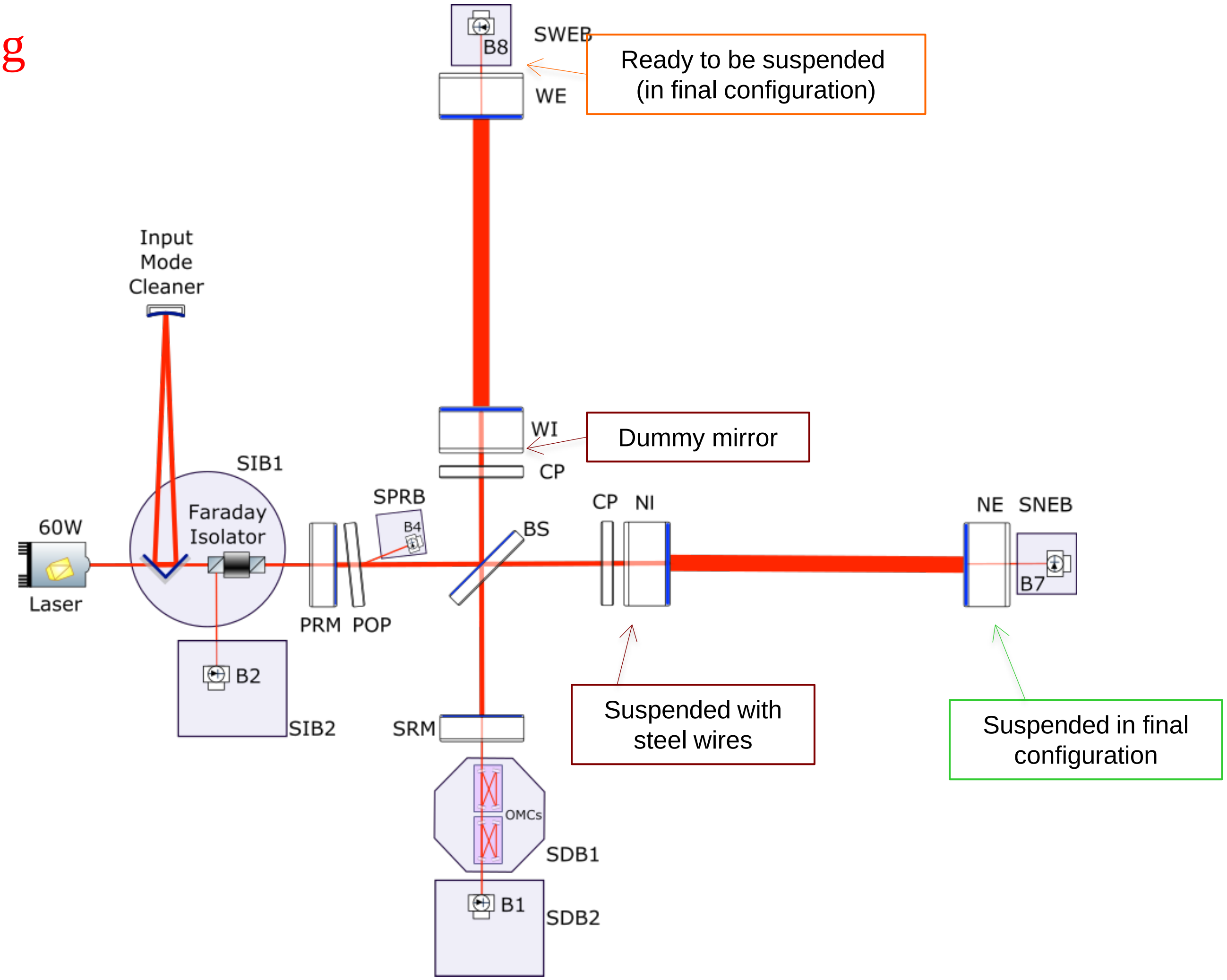
- Integration phase nearing completion
  - A few months delay due to two main issues
    - → 13 (out of ~300) superattenuator blades found broken
    - → 3 monolithic suspension failures after a few days under vacuum

- Broken blades
  - Origin of the problem found
  - Risky blades (40%) identified and replaced preventively
    - → Superattenuator completion delayed by a few months
  - Additional spare production
  - Procedure defined for fast in-situ replacements

- Monolithic suspension failures
  - Likely due to a production issue in a bunch of silica anchors
  - New (more robust) anchor design
  - New procedure defined to evacuate the towers
    - → One monolithic payload under vacuum for more than a month
    - → One mirror suspended with metal wires; two others not suspended yet

Advanced Virgo status

- What is currently missing
  - All the other mirrors in place for months



Ready to be suspended (in final configuration)

Dummy mirror

Suspended with steel wires

Suspended in final configuration

- Still some less crucial equipments to be installed
  - Parallel to the commissioning activities

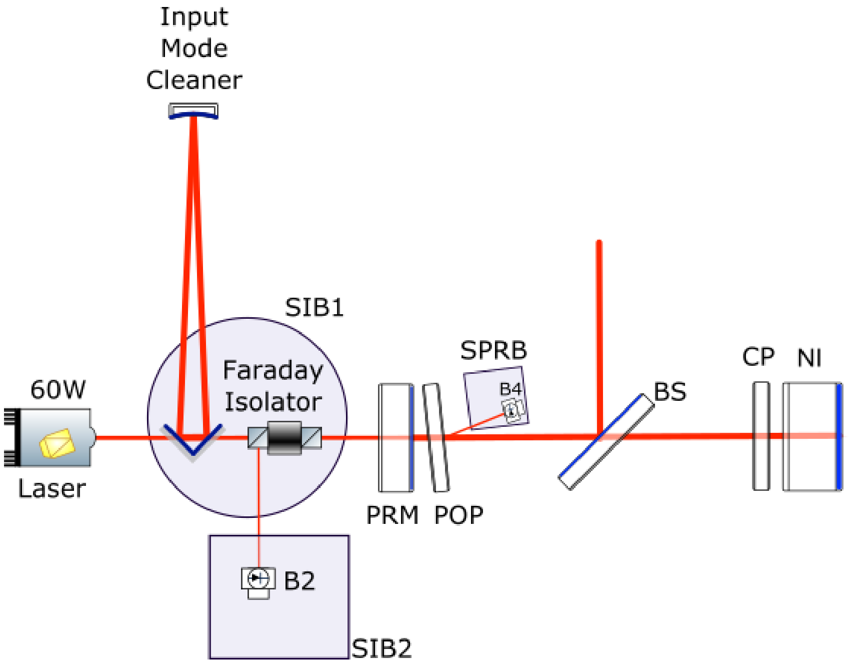
Advanced Virgo status

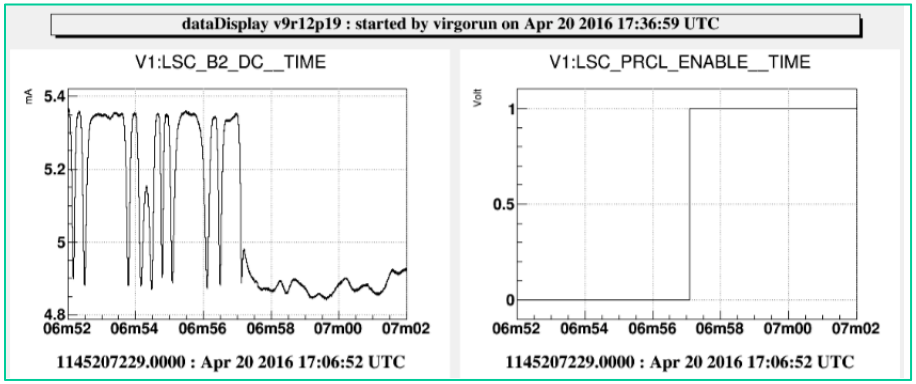
- All towers closed in the central building since last month

- All detection benches installed
- All cryotraps cooled down
- Commissioning of the injection system completed

Advanced Virgo status

- First lock of a cavity: power recycling → north input mirror





dataDisplay v9r12p19 : started by virgorun on Apr 20 2016 17:36:59 UTC

1145207229.0000 : Apr 20 2016 17:06:52 UTC

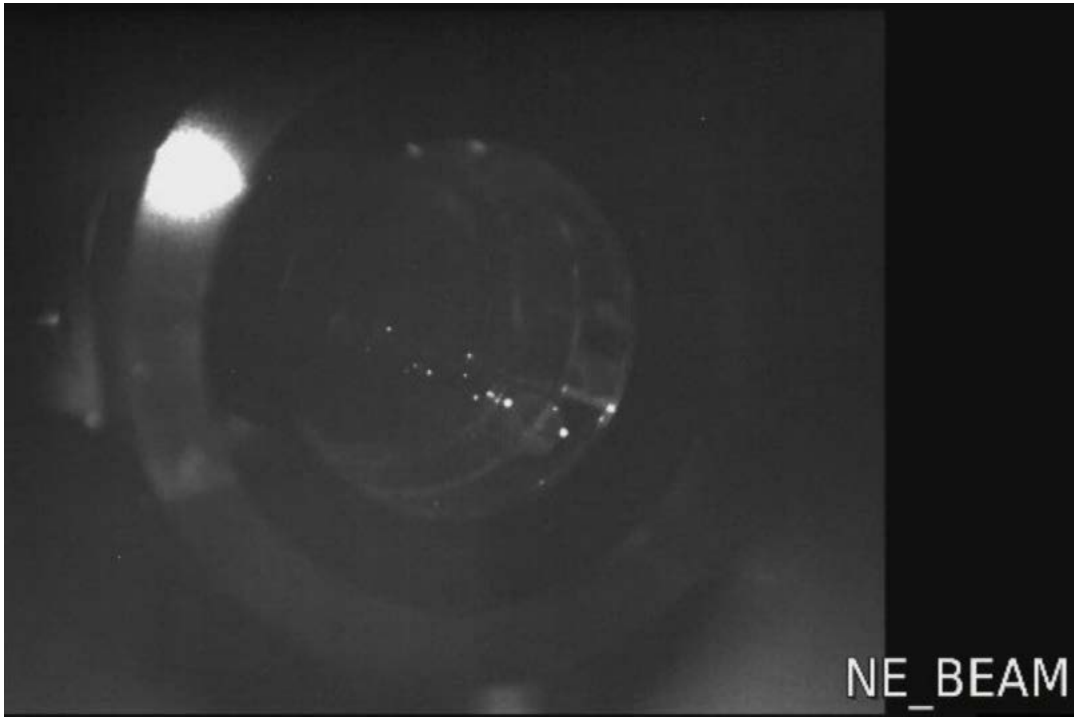
1145207229.0000 : Apr 20 2016 17:06:52 UTC

- Sensitivity: only 8 orders of magnitude to go…
- But: cavity locked with upgraded superattenuators, new payload design,
  new control electronics, digital demodulation,
  new acquisition/locking software, use of ring heater…

→ Nice integration test!

Advanced Virgo status

- Seeing the (laser) light at the end of the (3-km long) tunnel(s)!?



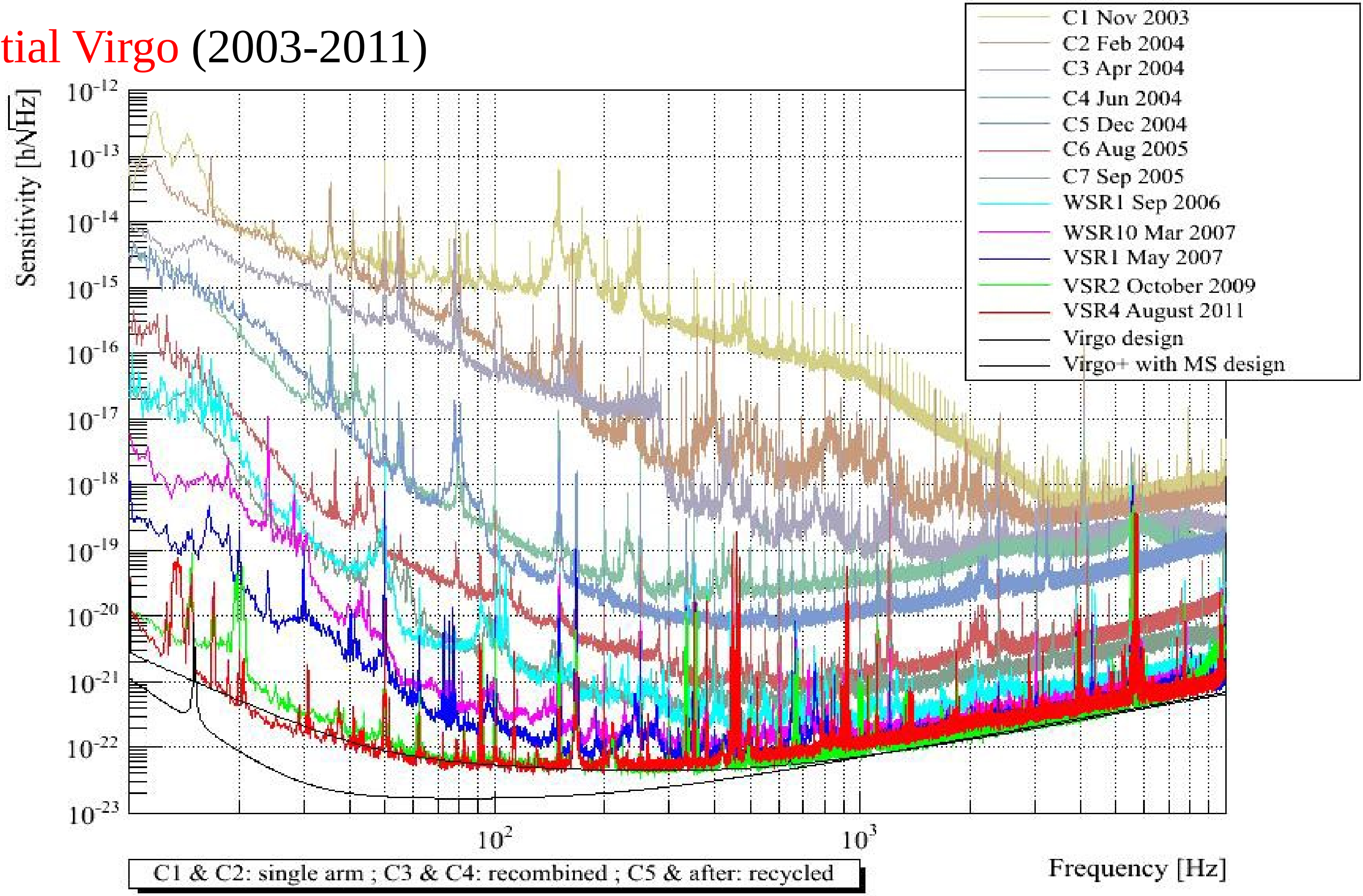
May 5: north end mirror payload hit by a direct beam coming from the injection system shortly after having opened the long arm vacuum valve

→ Transition from integration-dominated phase to commissioning

- Goal is still to join LIGO for the 2nd Observation Run (O2, end of 2016)

Improving the sensitivity: a long-term job

- Example of initial Virgo (2003-2011)



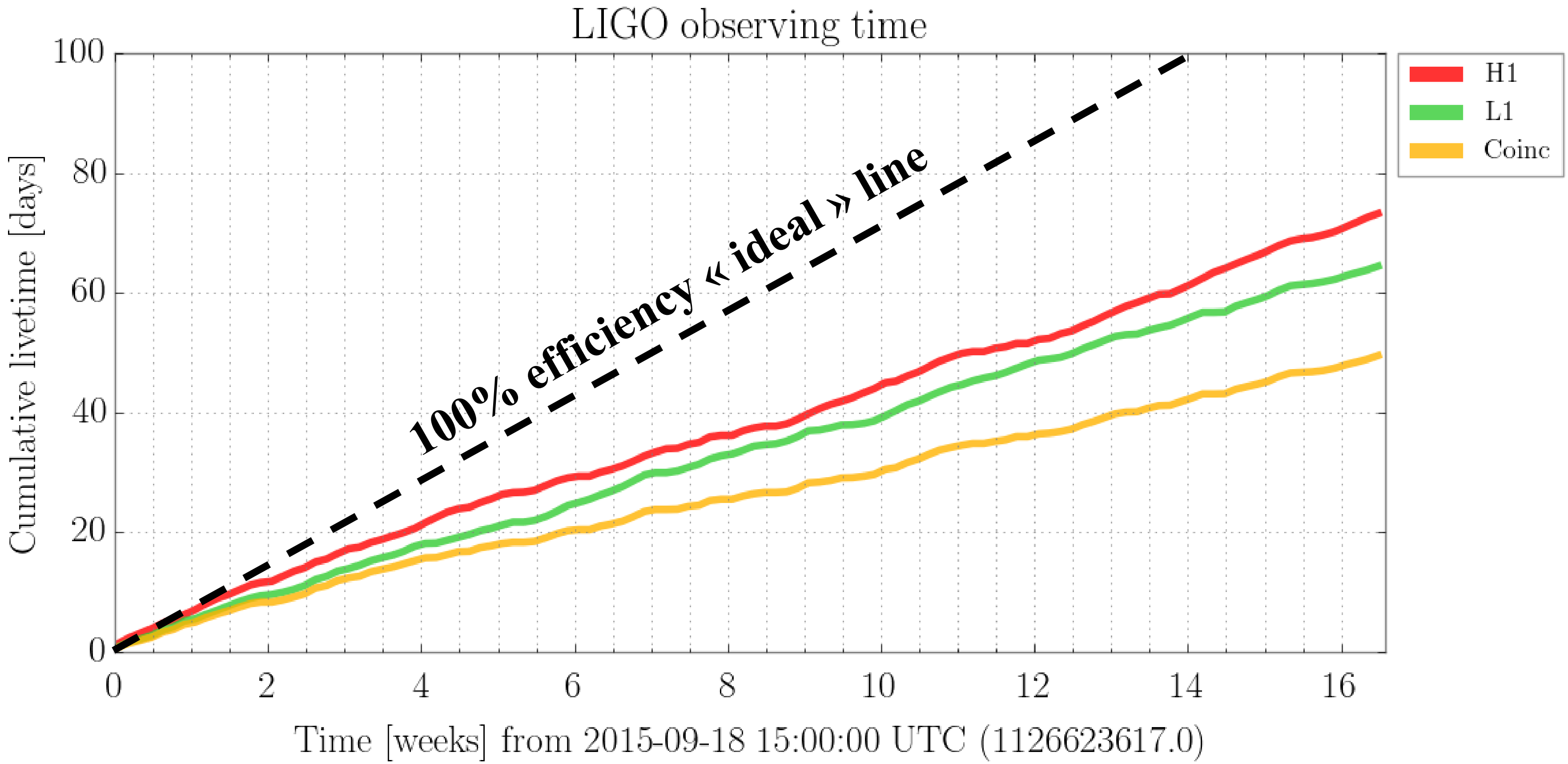
- Advanced LIGO detectors reached a record sensitivity much faster (< 1 year)
  - Experience gained and lessons learned from the first generation interferometers
  - Still room for improvement to reach the design sensitivity – and exceed it!

The Advanced LIGO «Observation 1» Run (2015/09 – 2016/01) & GW 150914

aLIGO O1 Run: Observing time

- September 2015 – January 2016
  - GW150914 showed up a few days before the official start of O1, during the « Engineering Run 8 »

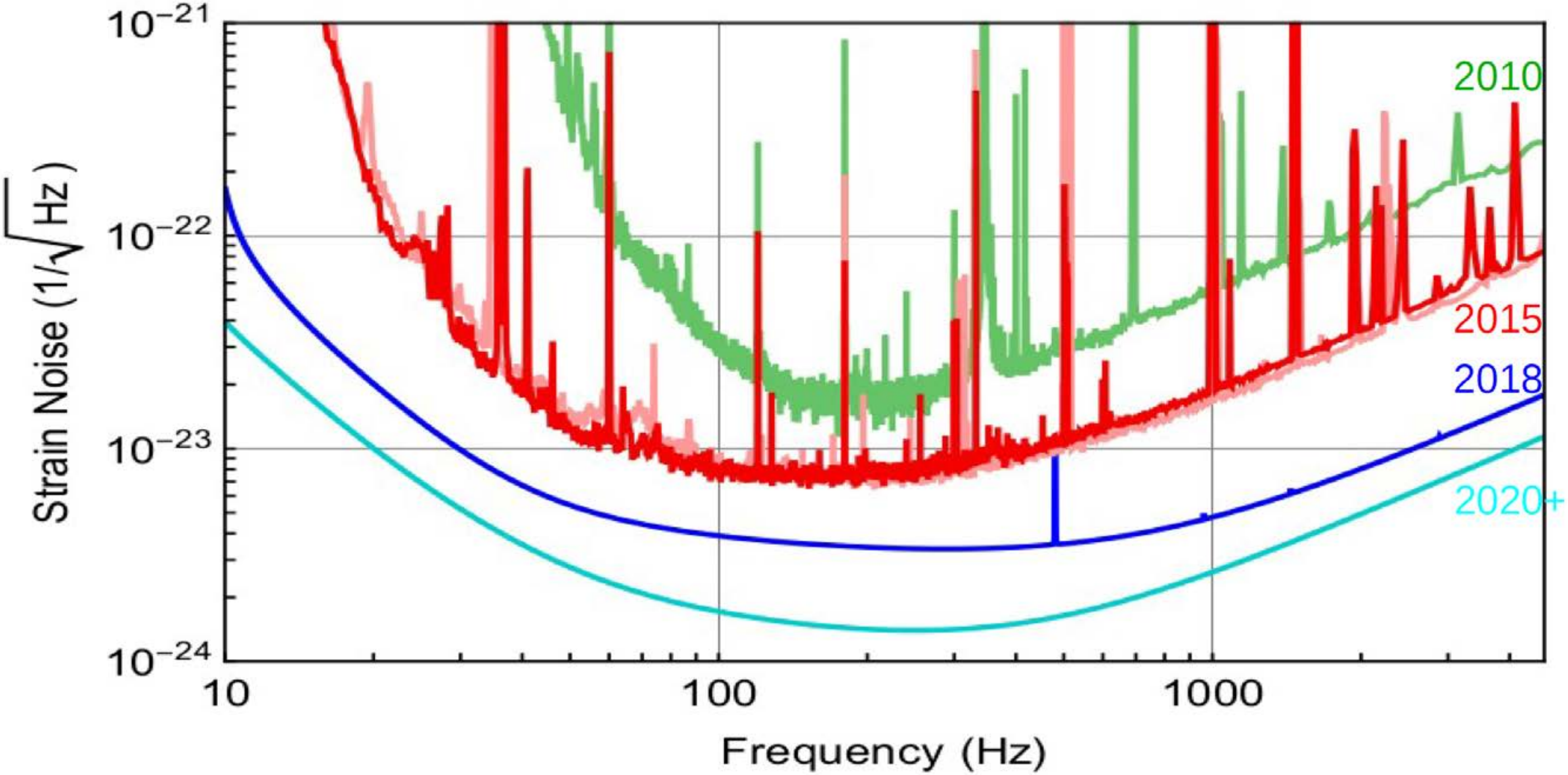
→ Both interferometers were already working nominally



aLIGO O1 Run: Sensitivity

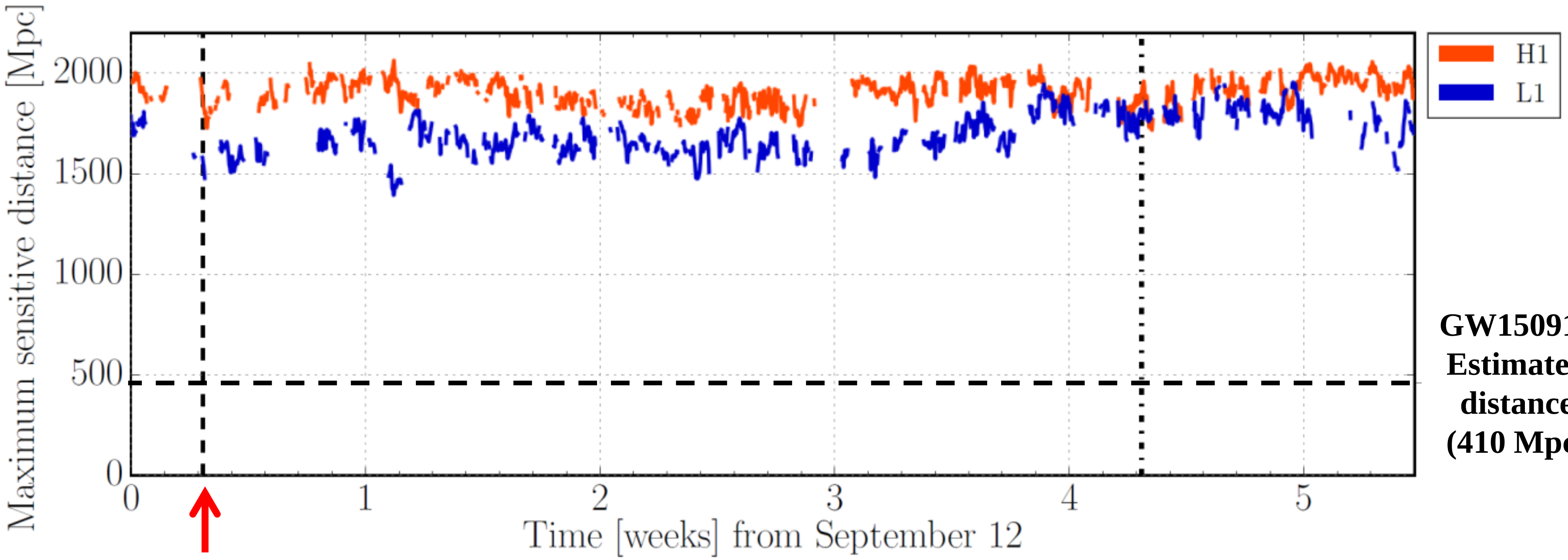
- Sensitiviy much improved with respect to the initial detectors
  - Factor 3-4 in strain

    → Factor 30-60 in volume probed
- Gain impressive at low frequency – where the signal GW150914 is located



aLIGO O1 Run: GW150914-like horizon

- Sky-averaged distance up to which a given signal can be detected
  - In this case a binary black hole system with the measured GW150914 parameters



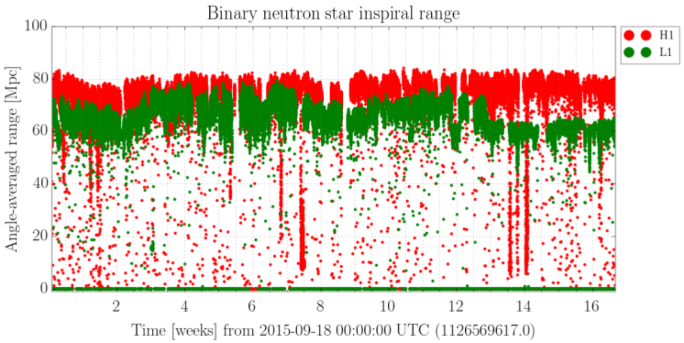
GW150914 Estimated distance (410 Mpc)

GW150914 time

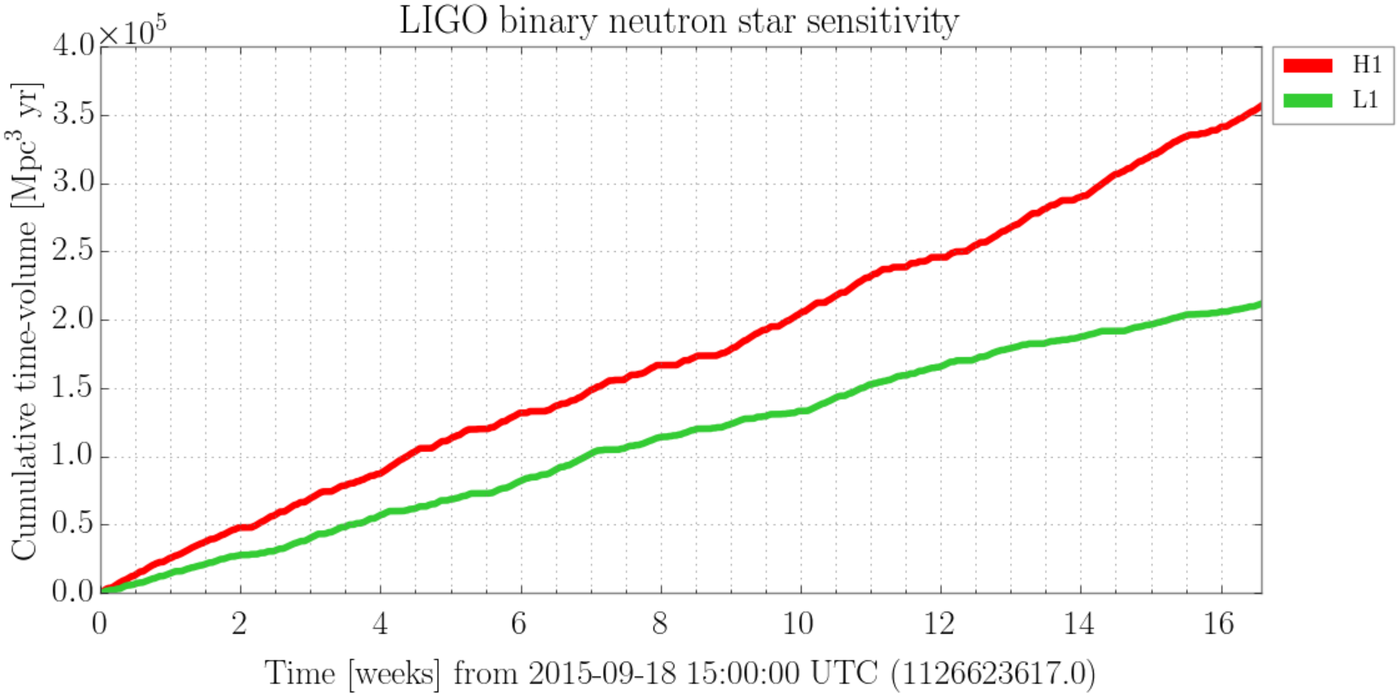
- Only depends on the actual sensitivity of the interferometer
  - Online monitoring tool used during data taking

aLIGO O1 Run: “VT” figure of merit

- Cumulative time-volume probed by the instruments
  → Expected number of sources (given a model)
  - Unit: Mpc^3.year
  - This slide: 1.4-1.4 $M_\odot$ « standard » binary neutron star system case



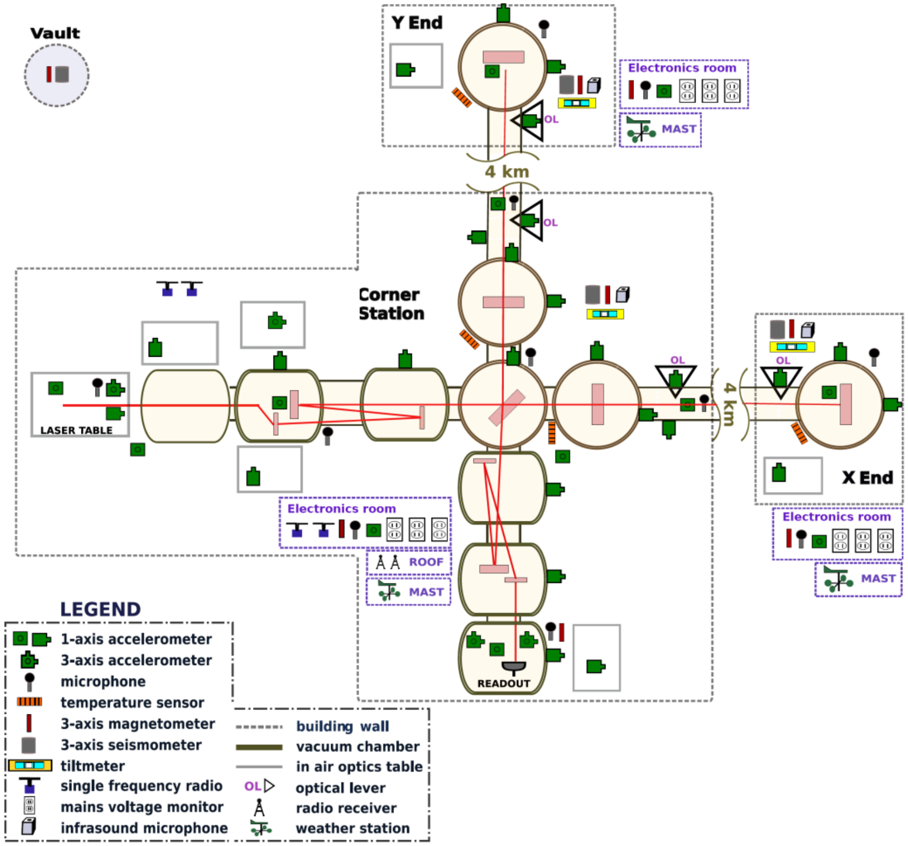
- Mixes sensitivity and duty cycle information



Data quality

- Detector configuration frozen to integrate enough data for background studies
  - ~40 days (until end of October) corresponding to 16 days of coincidence data

→ Steady performances over that period



- Tens of thousands of probes monitor the interferometer status and the environment
  - Virgo: h(t) ~ 100 kB/s
    DAQ ~ 30 MB/s

- Help identifying couplings with GW channel
  - Quantify how big a disturbance should be to produce such a large signal
  - Not to mention the distinctive shape of the GW150914 signal

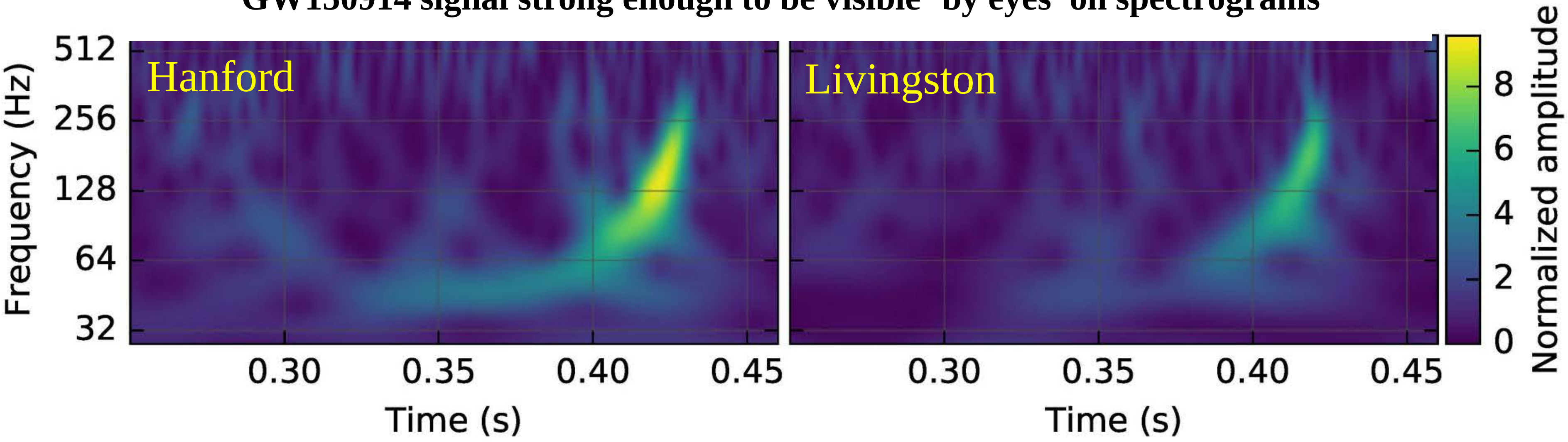
- Extensive studies performed
  - Uncorrelated and correlated noises
  - Bad data quality periods identified and vetoed

→ Clear conclusions: nominal running, no significant environmental disturbance

Burst search

- Search for clusters of excess power (above detector noise) in time-frequency plane
  - Wavelets

GW150914 signal strong enough to be visible ‘by eyes’ on spectrograms



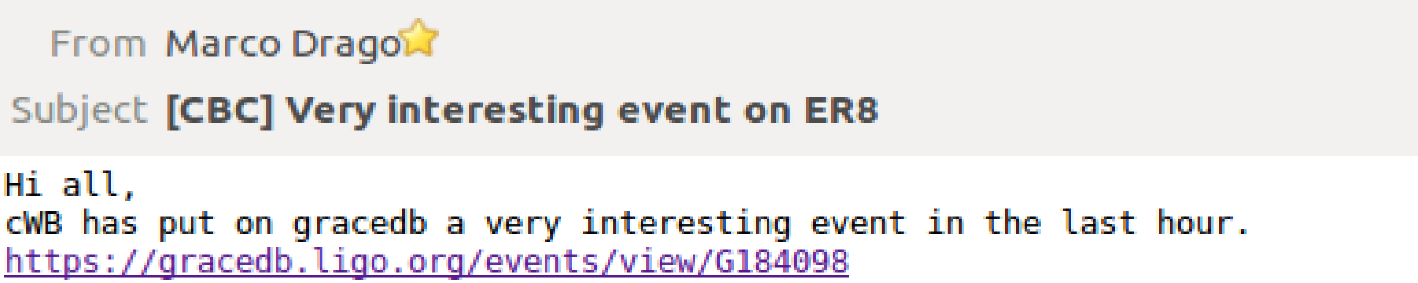
- Chirp-like shape: frequency and amplitude increasing with time

- Coherent excess in the two interferometers
  - Reconstructed signals required to be similar

- Efficiency similar to (optimal) matched filtering for binary black hole – short signal
  - Online last September for O1

Rapid response to GW150914

- 2015/09/14 11:51 CET: event recorded – first in Livingston, 7 ms later in Hanford
- 3 minutes later : event flagged, entry added to database, contacts notified
  - Online triggers important in particular for searches of counterparts
- 1 hour later: e-mails started flowing within the LIGO-Virgo collaboration



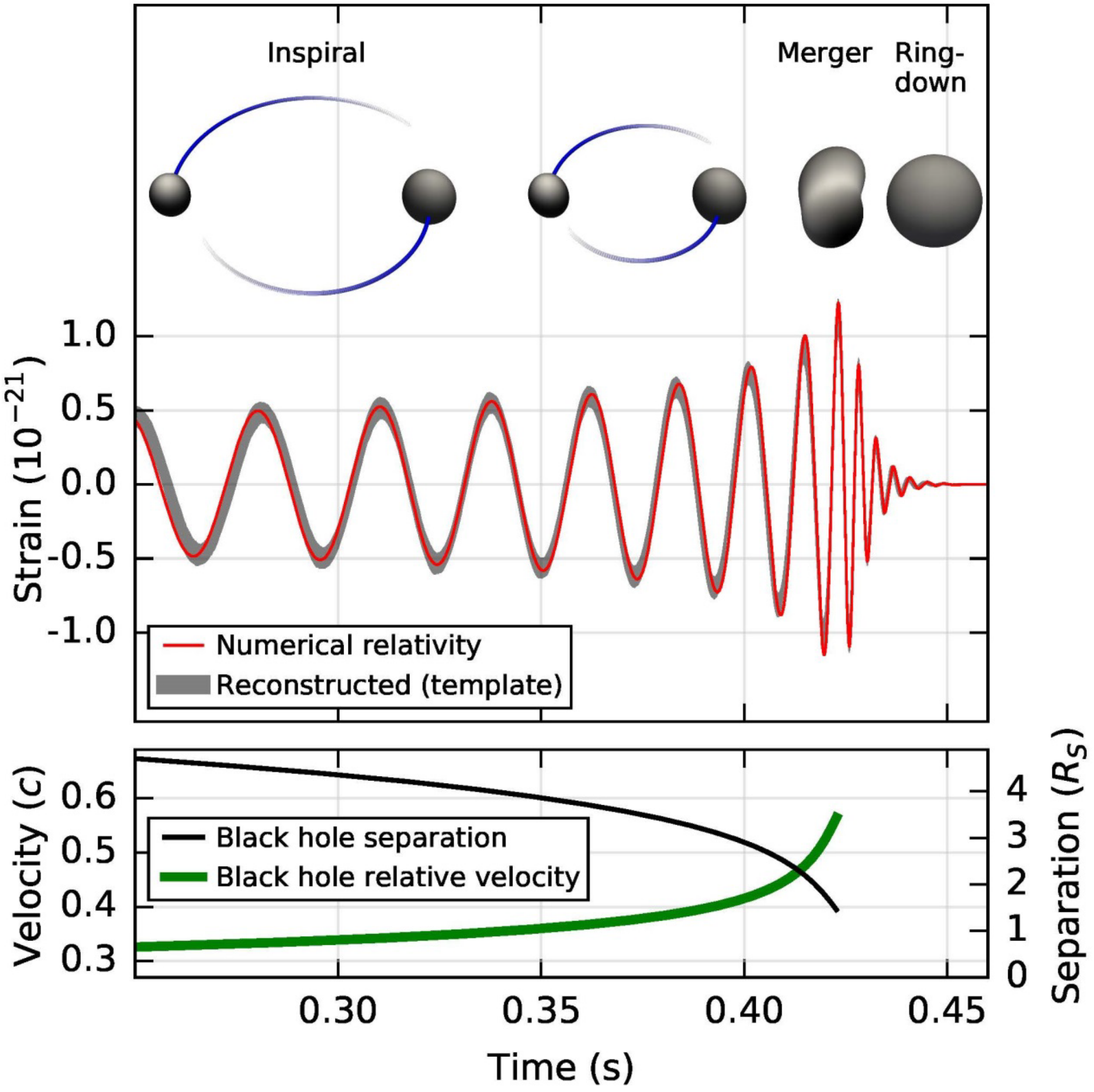
From Marco Drago

Subject **[CBC] Very interesting event on ER8**

```
Hi all,
cWB has put on gracedb a very interesting event in the last hour.
https://gracedb.ligo.org/events/view/G184098
```

- 20 minutes later: no signal injected at that time
  - Confirmed officially at 17:59 that day – blind injections useful to test pipelines
- 10 minutes later: binary black hole candidate
- 25 minutes later: data quality looks OK in both IFOs at the time of the event
- 15 minutes later: preliminary estimates of the signal parameters
  - False alarm rate < 1 / 300 years: a significant event!
- Two days later (09/16, 14:39 CET): alert circular sent to follow-up partners

Why two black holes?



- Result of matched filtering!
  - Excellent match between the best template and the measured signal
- Two massive compact objects orbiting around each other at 75 Hz (half the GW frequency), hence at relativistic speed, and getting very close before the merging: only a few R_S away!

→ Black holes are the only known objects which can fit this picture

- About 3 M_{Sun} radiated in GW
- The « brighest » event ever seen
  - More powerful than any gamma-ray burst detected so far
  - Peak power larger than 10 times the power emitted by the visible Universe

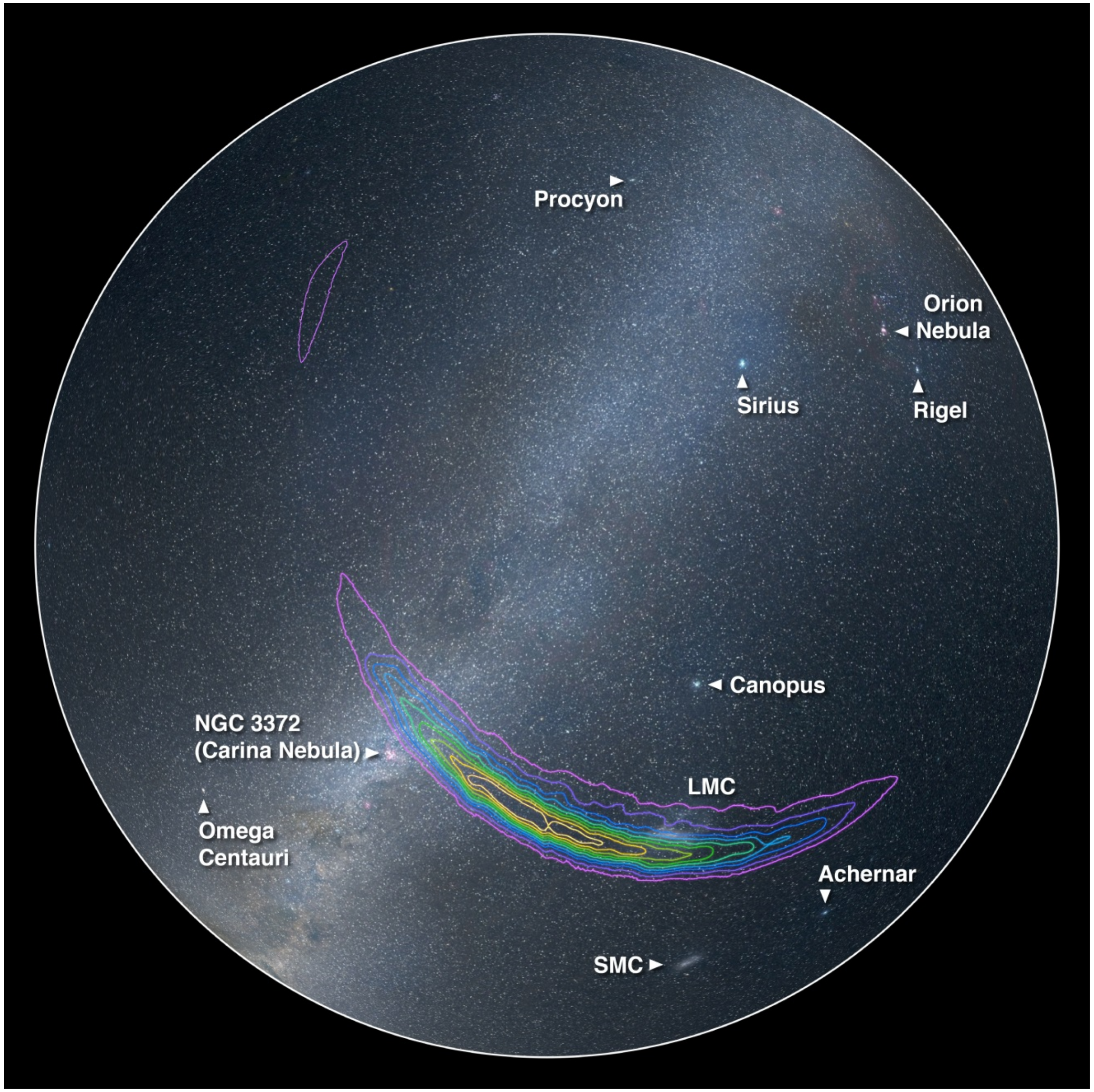
Skymap

- Sky at the time of the event

- Skymap contoured in deciles of probability

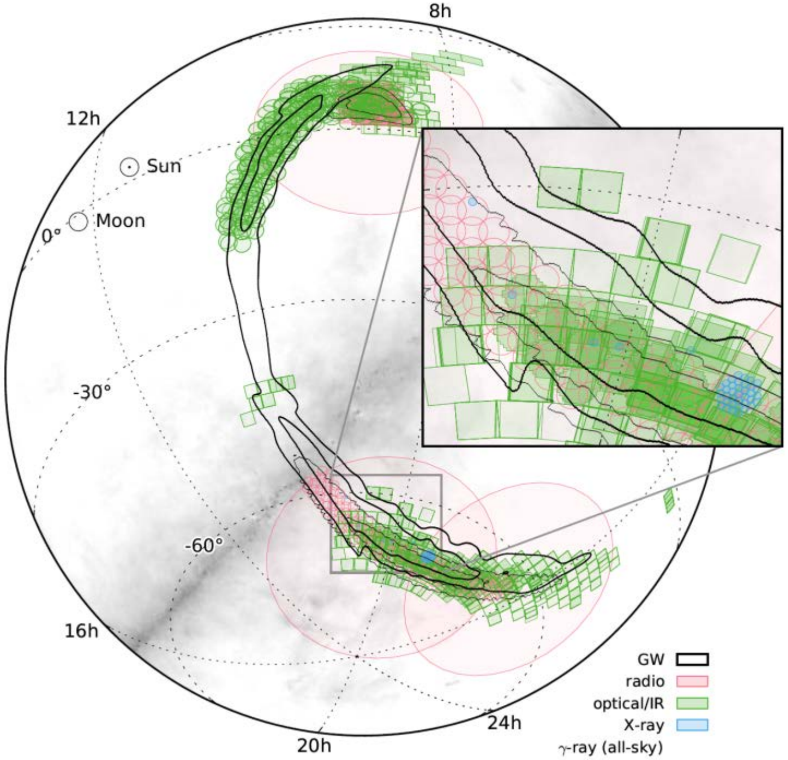
- 90% contour : ~ 590 degres^2

- View is from the South Atlantic Ocean, North at the top, with the Sun rising and the Milky Way diagonally from NW to SE

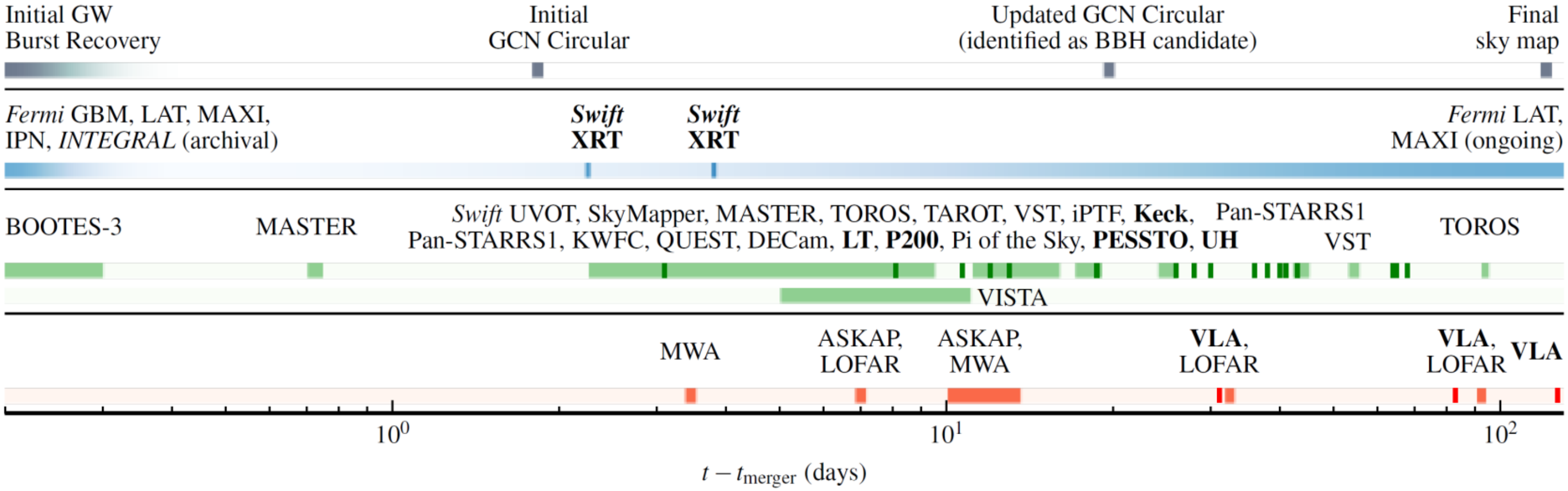


Looking for GW150914 counterparts

- Sky coverage



- Observation timeline: no counterpart found – none expected for a binary black hole



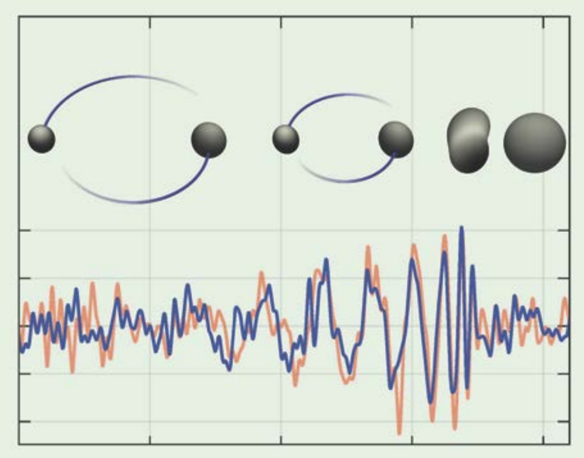
Conclusions

Outlook

- The network of advanced gravitational wave interferometers is taking shape
  - The two aLIGO detectors started taking data last September and detected the first direct gravitational wave signal (GW150914)
  - Virgo is completing its upgrade and is fully committed to joining LIGO asap
    → The right time for new groups to join the collaboration…
  - KAGRA should then join the network in 2018
  - And possibly a third LIGO detector (LIGO-India) some years later

- Sensitivity already good enough to detect gravitational waves
  - Improvements expected in the coming years
  - R&D activities already ongoing for 3rd generation instruments

- LIGO and Virgo will release results from the full « Observation 1 » run analysis in the coming weeks
  - Stay tuned…



GW detector peak sensitivity evolution vs. time