

Unreachable points in cosmological evolution in polynomial $f(R)$ gravity

References:

- JCAP 1602 (2016) no.02, 030
- JCAP 1410 (2014) no.10, 009

Collaborator: Saikat Chakraborty

Kaushik Bhattacharya

Indian Institute of Technology, Kanpur

Kanpur 208016, Uttar Pradesh

India.

Plan....

- $f(R)$ equations...
- Cosmological properties near the solutions of $f''(R) = 0$.
- Conclusion.

The $f(R)$ equations...

- In general instead of only the Ricci scalar, R , the Gravitational action has a general function $f(R)$
- Using FLRW metric $ds^2 = -dt^2 + a^2(t) [dr^2 + r^2(d\theta^2 + \sin^2 \theta d\phi^2)]$, the cosmological equations become,

$$3H^2 = \frac{\kappa}{F(R)} \rho_{\text{eff}},$$
$$3H^2 + 2\dot{H} = \frac{-\kappa}{F(R)} P_{\text{eff}},$$

where H is the conventional Hubble parameter. Here $F(R) \equiv \frac{df(R)}{dR}$.

The $f(R)$ equations...

- In general instead of only the Ricci scalar, R , the Gravitational action has a general function $f(R)$
- Using FLRW metric $ds^2 = -dt^2 + a^2(t) [dr^2 + r^2(d\theta^2 + \sin^2 \theta d\phi^2)]$, the cosmological equations become,

$$3H^2 = \frac{\kappa}{F(R)} \rho_{\text{eff}},$$
$$3H^2 + 2\dot{H} = \frac{-\kappa}{F(R)} P_{\text{eff}},$$

where H is the conventional Hubble parameter. Here $F(R) \equiv \frac{df(R)}{dR}$.

- Here

$$\rho_{\text{eff}} \equiv \rho + \rho_{\text{curv}}, \quad P_{\text{eff}} \equiv P + P_{\text{curv}},$$

where ρ and P specifies the properties of a perfect hydrodynamic fluid where $P = \omega\rho$.

The $f(R)$ equations...

- In general instead of only the Ricci scalar, R , the Gravitational action has a general function $f(R)$
- Using FLRW metric $ds^2 = -dt^2 + a^2(t) [dr^2 + r^2(d\theta^2 + \sin^2 \theta d\phi^2)]$, the cosmological equations become,

$$3H^2 = \frac{\kappa}{F(R)} \rho_{\text{eff}},$$

$$3H^2 + 2\dot{H} = \frac{-\kappa}{F(R)} P_{\text{eff}},$$

where H is the conventional Hubble parameter. Here $F(R) \equiv \frac{df(R)}{dR}$.

- Here

$$\rho_{\text{eff}} \equiv \rho + \rho_{\text{curv}}, \quad P_{\text{eff}} \equiv P + P_{\text{curv}},$$

where ρ and P specifies the properties of a perfect hydrodynamic fluid where $P = \omega\rho$.

- The curvature induced ρ_{curv} and P_{curv} are:

$$\rho_{\text{curv}} \equiv \frac{RF - f}{2\kappa} - \frac{3H\dot{R}F'}{\kappa}, \quad P_{\text{curv}} \equiv \frac{\dot{R}^2 F'' + 2H\dot{R}F' + \ddot{R}F'}{\kappa} - \frac{RF - f}{2\kappa}.$$

Properties of $f(R)$ theories..

- The field equations show that

$$F(R) \equiv f'(R) > 0,$$

for the **effective gravitational coupling**, given by $\frac{\kappa}{F(R)}$, to remain positive.

Properties of $f(R)$ theories..

- The field equations show that

$$F(R) \equiv f'(R) > 0 ,$$

for the **effective gravitational coupling**, given by $\frac{\kappa}{F(R)}$, to remain positive.

- The equations also show that if

$$F'(R) \equiv f''(R) > 0 ,$$

the theory is stable. If the condition is not specified then the **effective gravitational coupling** can increase with an increase of R and produce a **run-away effect**.

Properties of $f(R)$ theories..

- The field equations show that

$$F(R) \equiv f'(R) > 0 ,$$

for the **effective gravitational coupling**, given by $\frac{\kappa}{F(R)}$, to remain positive.

- The equations also show that if

$$F'(R) \equiv f''(R) > 0 ,$$

the theory is stable. If the condition is not specified then the **effective gravitational coupling** can increase with an increase of R and produce a **run-away effect**.

- For small scalar curvature and weak gravity, the above run-away effect is a real instability, called **Dolgov-Kawasaki instability**. For high values of R and strong gravity regime the instability may not be that dangerous.

Properties of $f(R)$ theories..

- The field equations show that

$$F(R) \equiv f'(R) > 0 ,$$

for the **effective gravitational coupling**, given by $\frac{\kappa}{F(R)}$, to remain positive.

- The equations also show that if

$$F'(R) \equiv f''(R) > 0 ,$$

the theory is stable. If the condition is not specified then the **effective gravitational coupling** can increase with an increase of R and produce a **run-away effect**.

- For small scalar curvature and weak gravity, the above run-away effect is a real instability, called **Dolgov-Kawasaki instability**. For high values of R and strong gravity regime the instability may not be that dangerous.
- The talk will address the cosmological phase when

$$F'(R = R_c) = 0 .$$

These particular phase shows interesting behavior in $f(R)$ theories.

Polynomial $f(R)$ theories and their properties

Out of the various forms of possible $f(R)$ the polynomial form remains a simple choice.

- One of the most interesting things about polynomial theories is, “One cannot simultaneously have both $F(R) > 0$ and $F'(R) > 0$ for all possible values of R in a polynomial $f(R)$ theory of gravity”. This is because,

Polynomial $f(R)$ theories and their properties

Out of the various forms of possible $f(R)$ the polynomial form remains a simple choice.

- One of the most interesting things about polynomial theories is, “One cannot simultaneously have both $F(R) > 0$ and $F'(R) > 0$ for all possible values of R in a polynomial $f(R)$ theory of gravity”. This is because,
- if $f(R)$ is an even order polynomial, then $F(R)$ must be an odd order polynomial and consequently it must have at least one real zero. As a result $F(R) > 0$ does not hold for all R .

Polynomial $f(R)$ theories and their properties

Out of the various forms of possible $f(R)$ the polynomial form remains a simple choice.

- One of the most interesting things about polynomial theories is, “One cannot simultaneously have both $F(R) > 0$ and $F'(R) > 0$ for all possible values of R in a polynomial $f(R)$ theory of gravity”. This is because,
- if $f(R)$ is an even order polynomial, then $F(R)$ must be an odd order polynomial and consequently it must have at least one real zero. As a result $F(R) > 0$ does not hold for all R .
- On the other hand suppose $f(R)$ is such an odd order polynomial so that the even order polynomial $F(R)$ has no real roots and the condition $F(R) > 0$ holds for all R . But then $F'(R)$ is an odd order polynomial and has at least one real root and consequently one cannot have $F'(R) > 0$ for all R .

Polynomial $f(R)$ theories and their properties

Out of the various forms of possible $f(R)$ the polynomial form remains a simple choice.

- One of the most interesting things about polynomial theories is, “One cannot simultaneously have both $F(R) > 0$ and $F'(R) > 0$ for all possible values of R in a polynomial $f(R)$ theory of gravity”. This is because,
- if $f(R)$ is an even order polynomial, then $F(R)$ must be an odd order polynomial and consequently it must have at least one real zero. As a result $F(R) > 0$ does not hold for all R .
- On the other hand suppose $f(R)$ is such an odd order polynomial so that the even order polynomial $F(R)$ has no real roots and the condition $F(R) > 0$ holds for all R . But then $F'(R)$ is an odd order polynomial and has at least one real root and consequently one cannot have $F'(R) > 0$ for all R .
- We will see that the value of R_c , for which $F'(R = R_c) = 0$, separates one region of cosmological existence from another different region.

The variables near $R = R_c$

- The field equations near $R = R_c$ where $F'(R_c) = 0$ can be written once we specify the values of the cosmological parameters to be H_c, ρ_c at that epoch.

The variables near $R = R_c$

- The field equations near $R = R_c$ where $F'(R_c) = 0$ can be written once we specify the values of the cosmological parameters to be H_c, ρ_c at that epoch.
- Near the (R_c, H_c, ρ_c) point

$$R = R_c + \delta R, \quad H = H_c + \delta H, \quad \rho = \rho_c + \delta \rho.$$

The variables near $R = R_c$

- The field equations near $R = R_c$ where $F'(R_c) = 0$ can be written once we specify the values of the cosmological parameters to be H_c, ρ_c at that epoch.
- Near the (R_c, H_c, ρ_c) point

$$R = R_c + \delta R, \quad H = H_c + \delta H, \quad \rho = \rho_c + \delta \rho.$$

- To linear order in fluctuations one can

$$f(R) = f(R_c) + F(R_c) \delta R,$$

and

$$F(R) = F(R_c), \quad F'(R) = F''(R_c) \delta R, \quad F''(R) = F''(R_c) + F'''(R_c) \delta R,$$

where we have assumed $F(R) > 0$ in general and $F'(R_c) = 0$.

The variables near $R = R_c$

- The field equations near $R = R_c$ where $F'(R_c) = 0$ can be written once we specify the values of the cosmological parameters to be H_c, ρ_c at that epoch.
- Near the (R_c, H_c, ρ_c) point

$$R = R_c + \delta R, \quad H = H_c + \delta H, \quad \rho = \rho_c + \delta \rho.$$

- To linear order in fluctuations one can

$$f(R) = f(R_c) + F(R_c) \delta R,$$

and

$$F(R) = F(R_c), \quad F'(R) = F''(R_c) \delta R, \quad F''(R) = F''(R_c) + F'''(R_c) \delta R,$$

where we have assumed $F(R) > 0$ in general and $F'(R_c) = 0$.

- Using this assumptions the field equations, **up to first order of the fluctuations**, near $R = R_c$ are.....

The field equations near $R = R_c$

$$\begin{aligned}3H_c^2 &= \frac{\kappa}{F(R_c)} \left[\rho_c + \frac{1}{2\kappa} \{R_c F(R_c) - f(R_c)\} \right], \\H_c \delta H &= \frac{\kappa}{6F(R_c)} \delta \rho, \\ \dot{H} &= \delta \dot{H} = -\frac{\kappa(1+\omega)}{2F(R_c)} \delta \rho,\end{aligned}$$

and $\rho_c(1+\omega) = 0$, where ω specifies the hydrodynamic fluid.

The field equations near $R = R_c$

$$\begin{aligned}3H_c^2 &= \frac{\kappa}{F(R_c)} \left[\rho_c + \frac{1}{2\kappa} \{R_c F(R_c) - f(R_c)\} \right], \\H_c \delta H &= \frac{\kappa}{6F(R_c)} \delta\rho, \\ \dot{H} &= \delta\dot{H} = -\frac{\kappa(1+\omega)}{2F(R_c)} \delta\rho,\end{aligned}$$

and $\rho_c(1+\omega) = 0$, where ω specifies the hydrodynamic fluid.

- If at $t = 0$ the system has $R = R_c$ then the universe must be **matter-less** or have a **fluid** with $\omega = -1$, if

$$\rho_c + \frac{1}{2\kappa} \{R_c F(R_c) - f(R_c)\} > 0.$$

The field equations near $R = R_c$

$$\begin{aligned}3H_c^2 &= \frac{\kappa}{F(R_c)} \left[\rho_c + \frac{1}{2\kappa} \{R_c F(R_c) - f(R_c)\} \right], \\H_c \delta H &= \frac{\kappa}{6F(R_c)} \delta \rho, \\ \dot{H} &= \delta \dot{H} = -\frac{\kappa(1+\omega)}{2F(R_c)} \delta \rho,\end{aligned}$$

and $\rho_c(1+\omega) = 0$, where ω specifies the hydrodynamic fluid.

- If at $t = 0$ the system has $R = R_c$ then the universe must be **matter-less** or have a **fluid** with $\omega = -1$, if

$$\rho_c + \frac{1}{2\kappa} \{R_c F(R_c) - f(R_c)\} > 0.$$

- If such a configuration exists, then in both the possible cases $\dot{H} = \delta \dot{H} = 0$. The universe is in a **perfect de Sitter phase** and will remain so for **infinite time**.

The field equations near $R = R_c$

$$\begin{aligned}3H_c^2 &= \frac{\kappa}{F(R_c)} \left[\rho_c + \frac{1}{2\kappa} \{R_c F(R_c) - f(R_c)\} \right], \\H_c \delta H &= \frac{\kappa}{6F(R_c)} \delta \rho, \\ \dot{H} &= \delta \dot{H} = -\frac{\kappa(1+\omega)}{2F(R_c)} \delta \rho,\end{aligned}$$

and $\rho_c(1+\omega) = 0$, where ω specifies the hydrodynamic fluid.

- If at $t = 0$ the system has $R = R_c$ then the universe must be **matter-less** or have a **fluid** with $\omega = -1$, if

$$\rho_c + \frac{1}{2\kappa} \{R_c F(R_c) - f(R_c)\} > 0.$$

- If such a configuration exists, then in both the possible cases $\dot{H} = \delta H = 0$. The universe is in a **perfect de Sitter phase** and will remain so for **infinite time**.
- If at $t = 0$ the system has $R \neq R_c$ while $\omega = -1$ (and $\rho = \text{constant}$), or $\rho = 0$ that initial state can reach the (R_c, H_c, ρ_c) point and get **stuck**, as it again hits a de Sitter like phase where $\dot{H} = 0$.

Cosmological behavior near $R = R_c$

- On the other hand if at $t = 0$ the system has $R \neq R_c$ while $\omega \neq -1$ and $\rho \neq 0$ then one can show that the system will require an infinite time to reach the (R_c, H_c, ρ_c) point. The reason is as follows.

Cosmological behavior near $R = R_c$

- On the other hand if at $t = 0$ the system has $R \neq R_c$ while $\omega \neq -1$ and $\rho \neq 0$ then one can show that the system will require an infinite time to reach the (R_c, H_c, ρ_c) point. The reason is as follows.
- Writing δH as h , the field equations reveal that near the (R_c, H_c, ρ_c) point

$$\dot{h} = -Ch,$$

for a constant C .

Cosmological behavior near $R = R_c$

- On the other hand if at $t = 0$ the system has $R \neq R_c$ while $\omega \neq -1$ and $\rho \neq 0$ then one can show that the system will require an infinite time to reach the (R_c, H_c, ρ_c) point. The reason is as follows.
- Writing δH as h , the field equations reveal that near the (R_c, H_c, ρ_c) point

$$\dot{h} = -Ch,$$

for a constant C .

The time to reach the (R_c, H_c, ρ_c) point is

$$t = \frac{1}{C} \ln \left[\frac{h}{\epsilon} \right].$$

As $\epsilon \rightarrow 0$ one must have $h \rightarrow 0$. But

$$t \rightarrow \infty, \text{ as } \epsilon \rightarrow 0.$$

Cosmological behavior near $R = R_c$

- On the other hand if at $t = 0$ the system has $R \neq R_c$ while $\omega \neq -1$ and $\rho \neq 0$ then one can show that the system will require an infinite time to reach the (R_c, H_c, ρ_c) point. The reason is as follows.
- Writing δH as h , the field equations reveal that near the (R_c, H_c, ρ_c) point

$$\dot{h} = -Ch,$$

for a constant C .

The time to reach the (R_c, H_c, ρ_c) point is

$$t = \frac{1}{C} \ln \left[\frac{h}{\epsilon} \right].$$

As $\epsilon \rightarrow 0$ one must have $h \rightarrow 0$. But

$$t \rightarrow \infty, \text{ as } \epsilon \rightarrow 0.$$

- Consequently, any arbitrary cosmological system will not be able to cross the (R_c, H_c, ρ_c) point in finite time.

Conclusion

- There are some configurations in cosmological space which are unreachable from other possible initial cosmological set up, when cosmological evolution is guided by $f(R)$ theories of gravity. Regions where $f''(R) > 0$ will be unreachable if initially one starts in an universe where $f''(R) < 0$. The Ricci scalar R_c for which $f''(R_c) = 0$ separates the reachable and unreachable points.

Conclusion

- There are some configurations in cosmological space which are unreachable from other possible initial cosmological set up, when cosmological evolution is guided by $f(R)$ theories of gravity. Regions where $f''(R) > 0$ will be unreachable if initially one starts in an universe where $f''(R) < 0$. The Ricci scalar R_c for which $f''(R_c) = 0$ separates the reachable and unreachable points.
- If somehow the universe gets trapped around this configuration where $f''(R_c) = 0$ then the universe will always remain in that configuration for eternity.

Conclusion

- There are some configurations in cosmological space which are unreachable from other possible initial cosmological set up, when cosmological evolution is guided by $f(R)$ theories of gravity. Regions where $f''(R) > 0$ will be unreachable if initially one starts in an universe where $f''(R) < 0$. The Ricci scalar R_c for which $f''(R_c) = 0$ separates the reachable and unreachable points.
- If somehow the universe gets trapped around this configuration where $f''(R_c) = 0$ then the universe will always remain in that configuration for eternity.
- Many theories of $f(R)$ gravity does not at all possess this kind of a configuration as $f(R) = R + \alpha R^2$ or $f(R) = \frac{1}{\alpha} [1 - \exp(\alpha R)]$. On the other hand most polynomial $f(R)$ theories whose order is cubic or more must have these kind of interesting configurations.

Conclusion

- There are some configurations in cosmological space which are unreachable from other possible initial cosmological set up, when cosmological evolution is guided by $f(R)$ theories of gravity. Regions where $f''(R) > 0$ will be unreachable if initially one starts in an universe where $f''(R) < 0$. The Ricci scalar R_c for which $f''(R_c) = 0$ separates the reachable and unreachable points.
- If somehow the universe gets trapped around this configuration where $f''(R_c) = 0$ then the universe will always remain in that configuration for eternity.
- Many theories of $f(R)$ gravity does not at all possess this kind of a configuration as $f(R) = R + \alpha R^2$ or $f(R) = \frac{1}{\alpha} [1 - \exp(\alpha R)]$. On the other hand most polynomial $f(R)$ theories whose order is cubic or more must have these kind of interesting configurations.
- The possible roots of the equation $f''(R) = 0$ partitions the probable regions of cosmological existence in cubic or higher order polynomial theories of gravity.