Signatures of Earth-Shadowing in the Direct Detection of Dark Matter

Bradley J. Kavanagh
LPTHE - Paris VI
with Riccardo Catena and Chris Kouvaris

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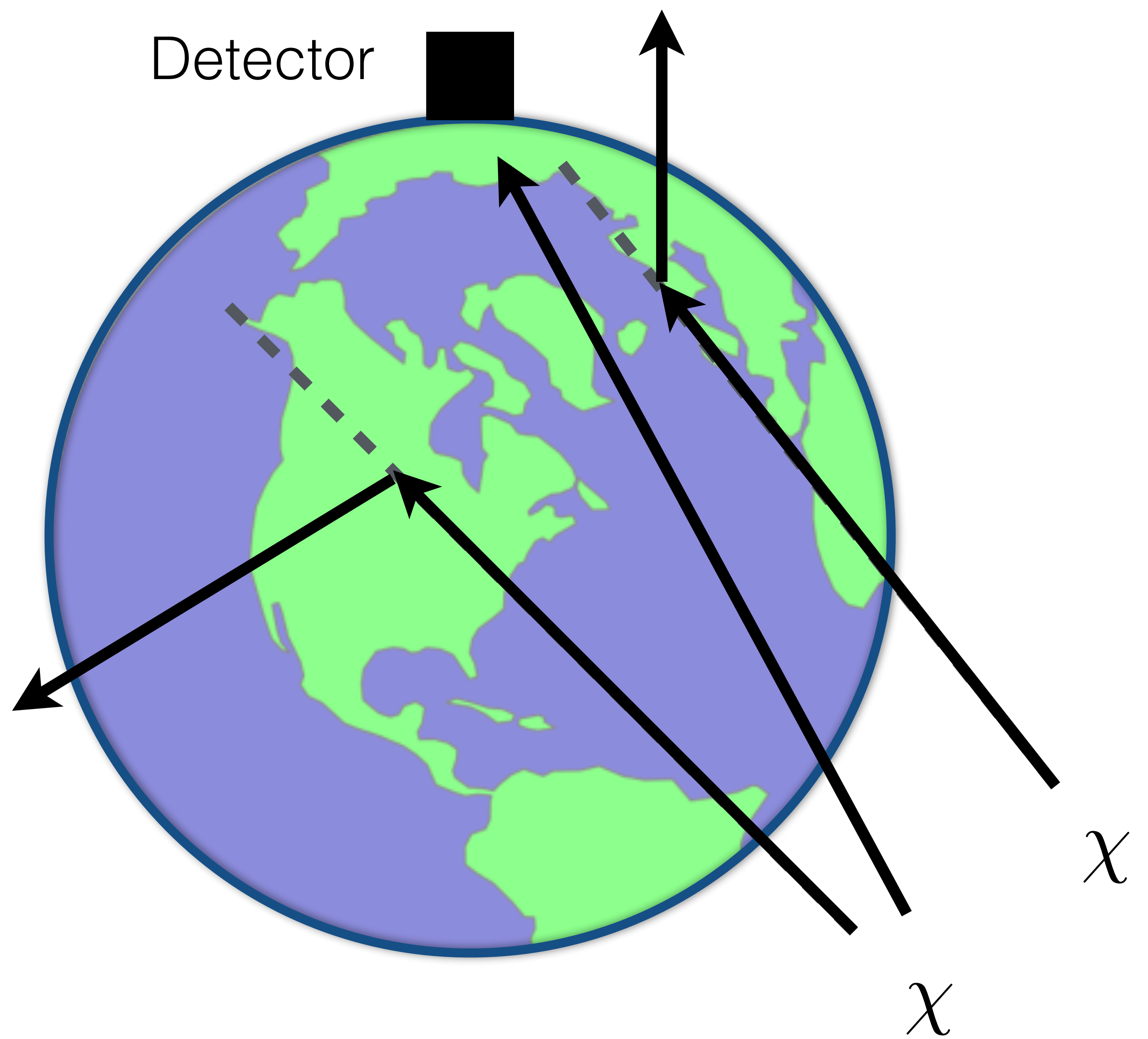
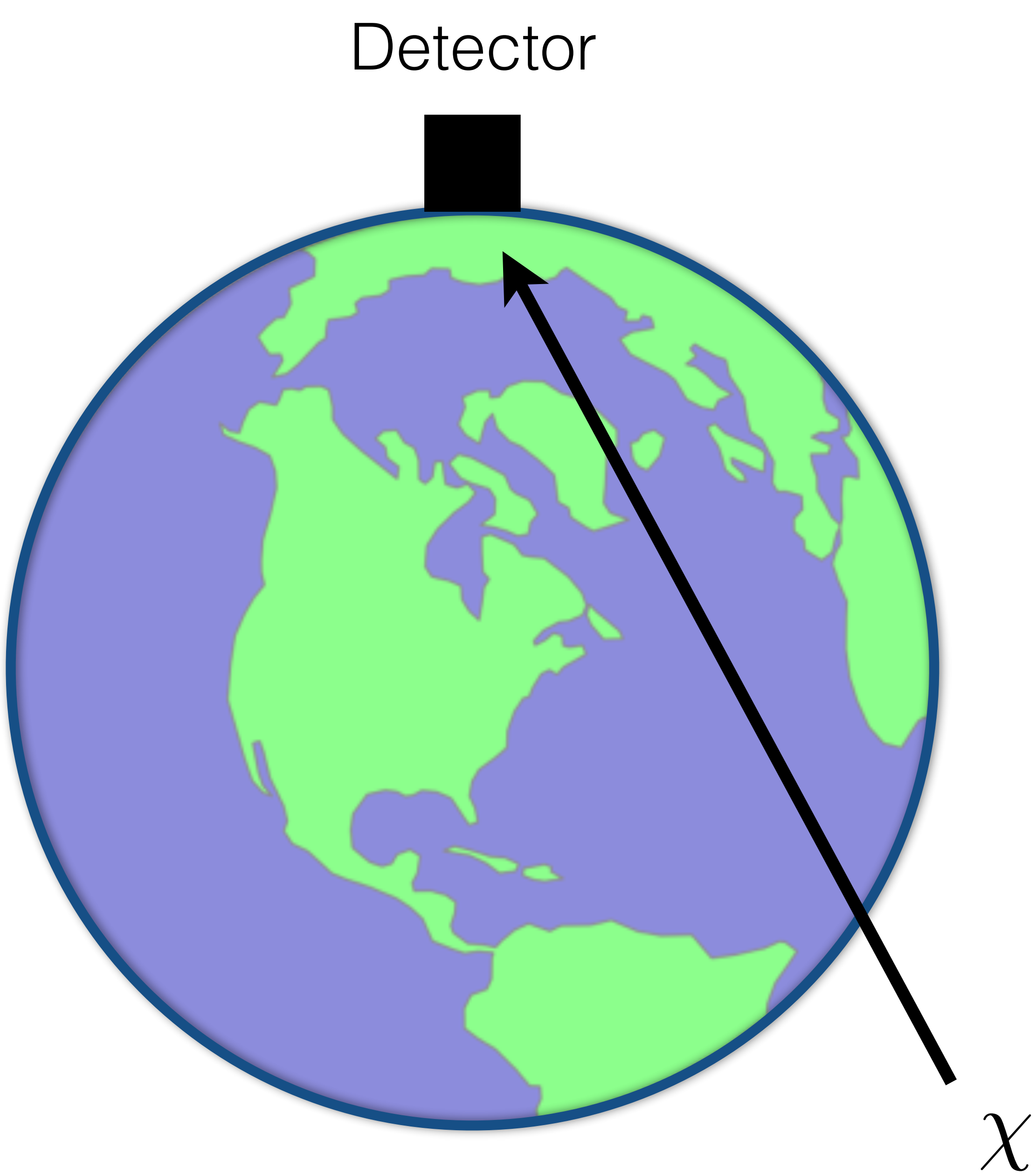
bkavanagh@lpthe.jussieu.fr
@BradleyKavanagh

Earth-Shadowing

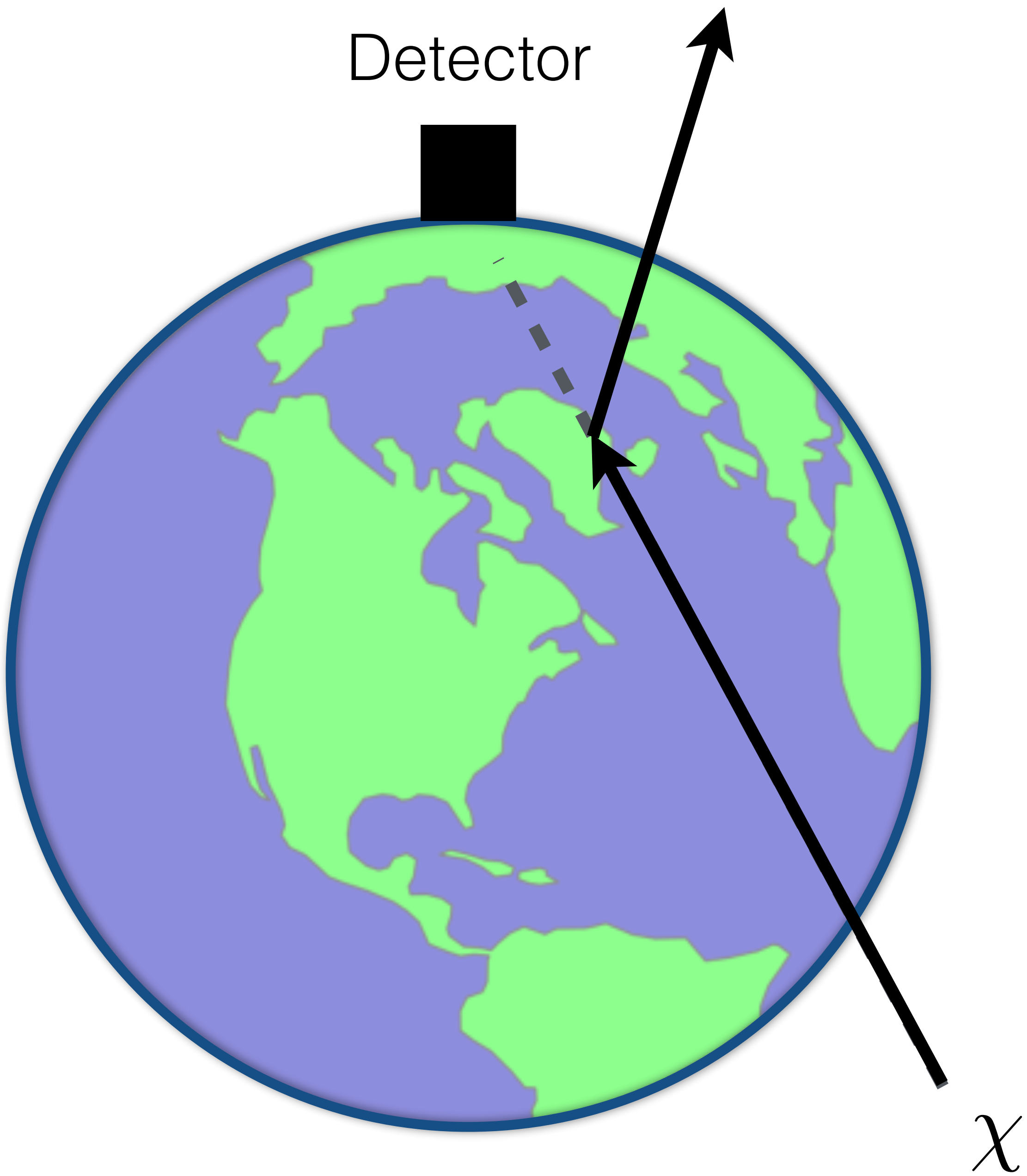


Earth-Shadowing



Unscattered (free) DM: $f_0(\mathbf{v})$

Earth-Shadowing - Attenuation



Previous calculations usually only consider DM attenuation

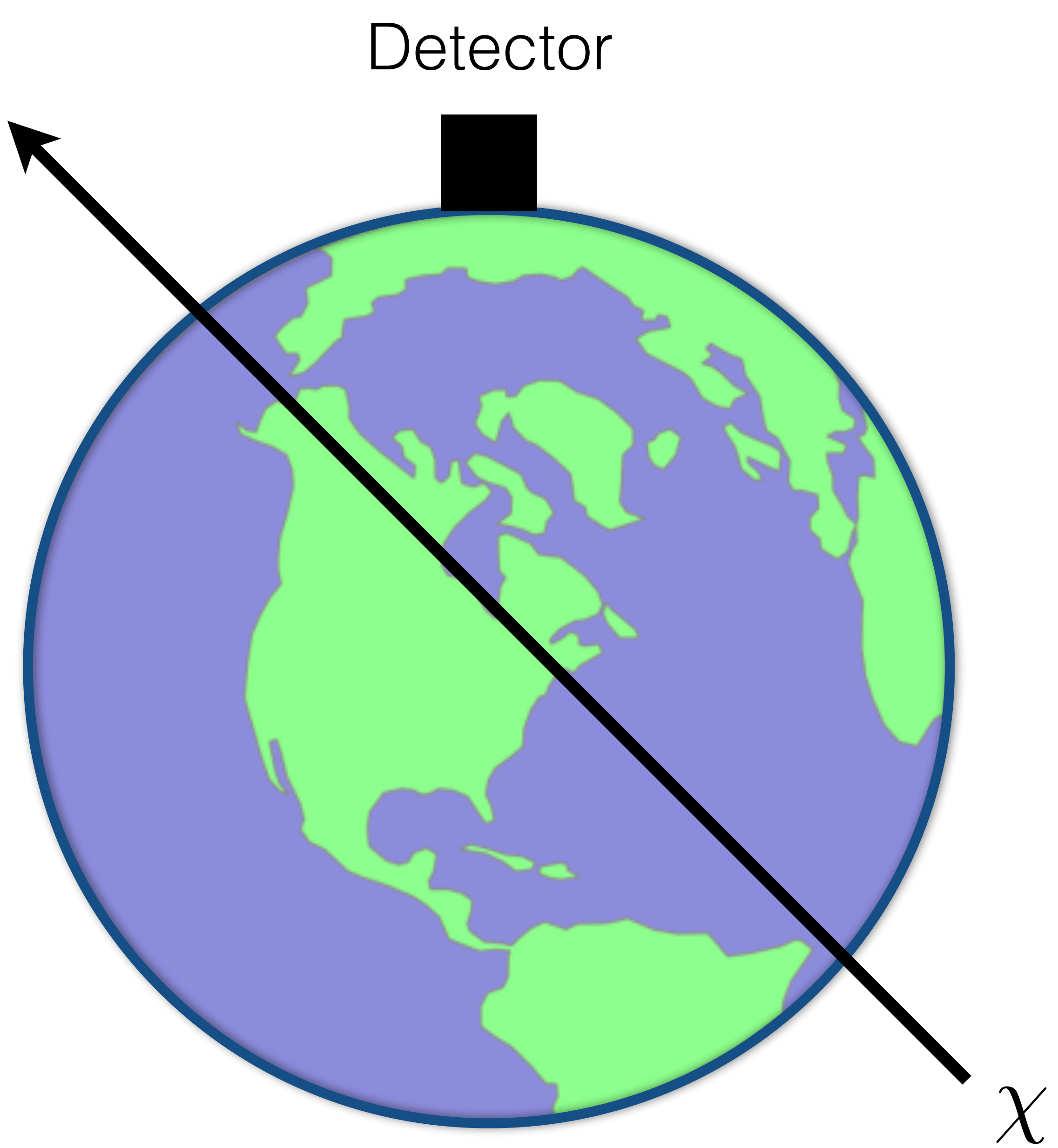
Zaharijas & Farrar [astro-ph/0406531]

Kouvaris & Shoemaker [1405.1729,1509.08720]

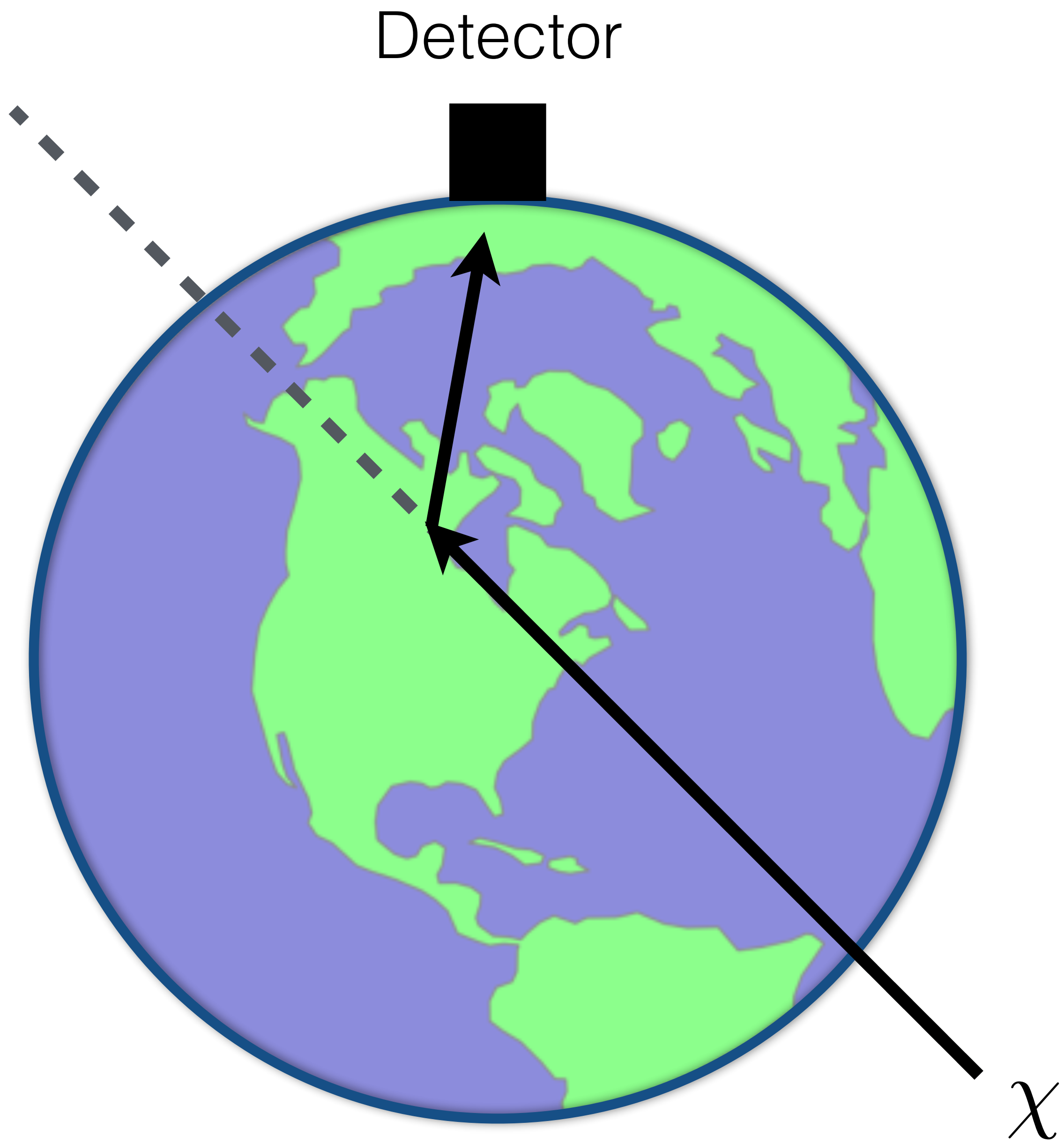
DAMA [1505.05336]

Attenuation of DM flux: $f(\mathbf{v}) \rightarrow f_0(\mathbf{v}) - f_A(\mathbf{v})$

Earth-Shadowing - Deflection



Earth-Shadowing - Deflection



Assuming DM mean free path

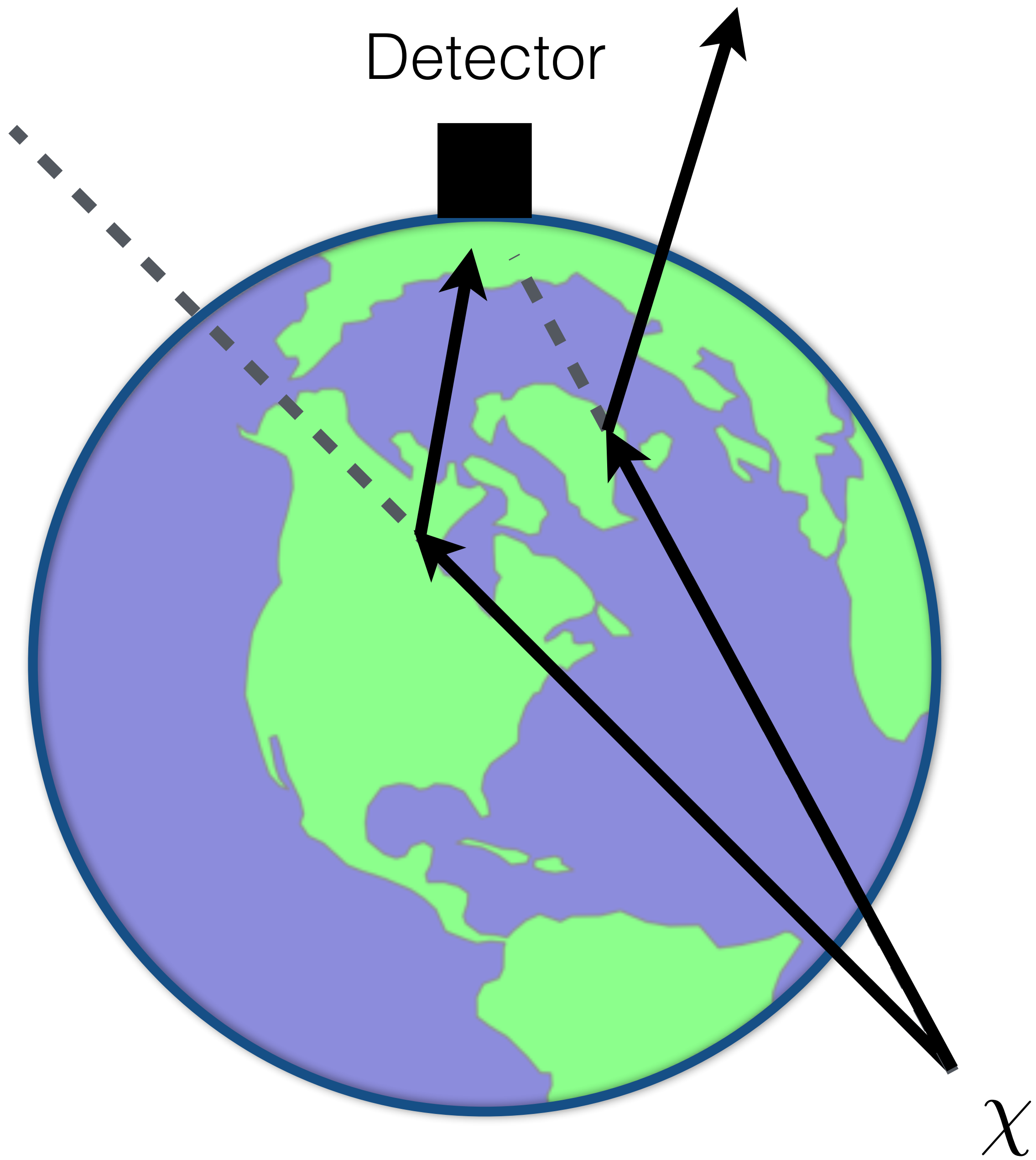
$$\lambda \gtrsim R_E$$

Considered in early Monte Carlo simulations

Collar & Avignone [PLB 275, 1992 and others]

We'll use the 'single scatter' approximation…

Earth-Shadowing



Assuming DM mean free path

$$\lambda \gtrsim R_E$$

Total DM velocity distribution: $f(\mathbf{v}) = f_0(\mathbf{v}) - f_A(\mathbf{v}) + f_D(\mathbf{v})$

→ altered flux, daily modulation, directionality…

Outline

Dark Matter (DM) and Direct Detection

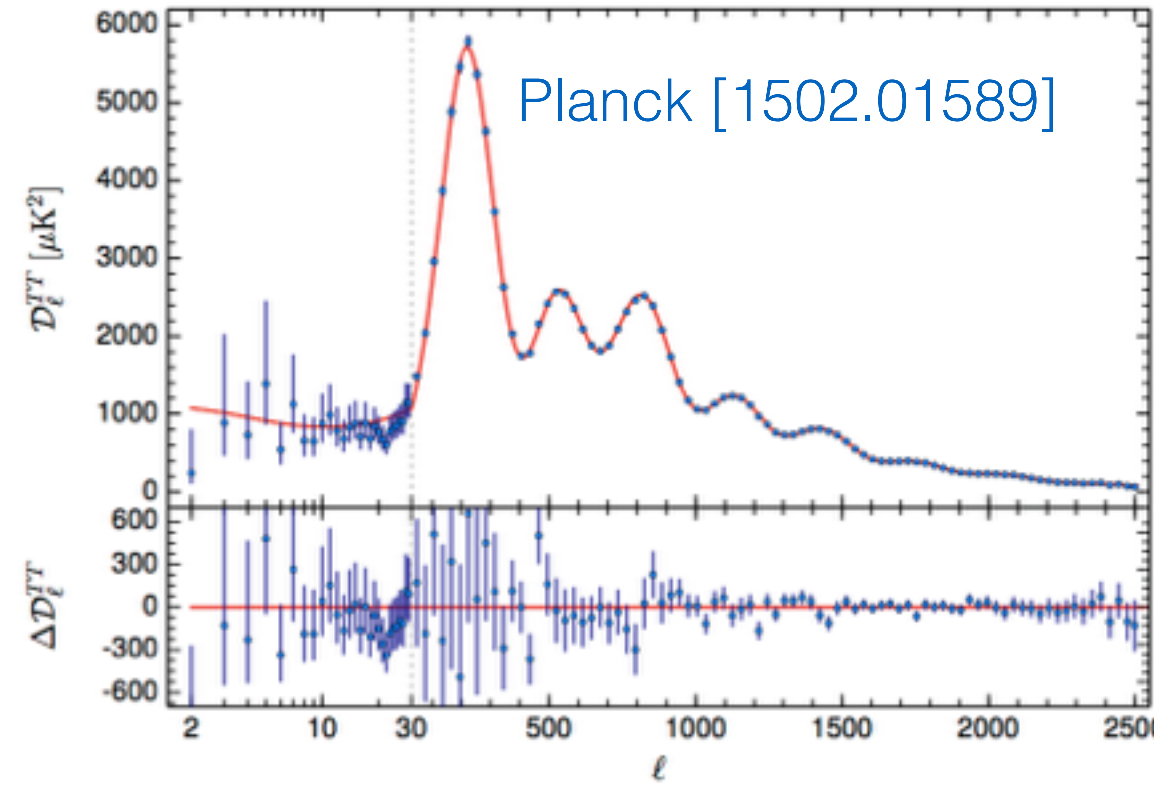
Calculating the Earth-Shadowing effect

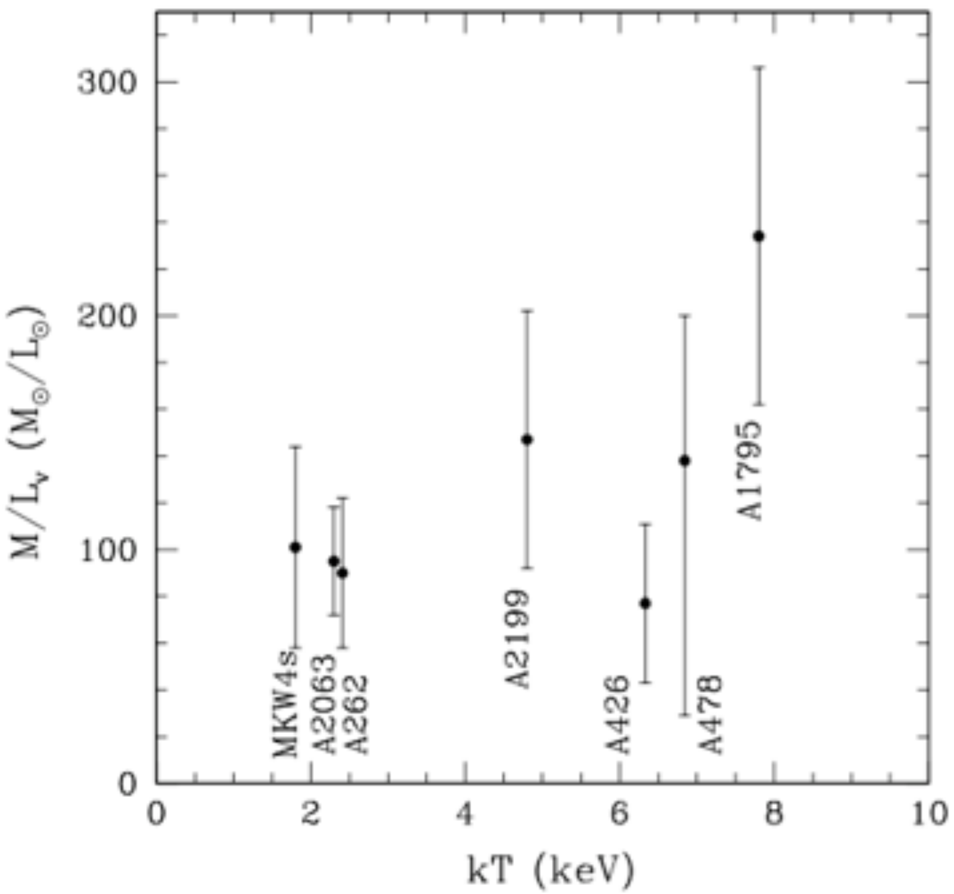
Non-relativistic Effective Field Theory of DM

Impact on the DM velocity distribution and modulation signatures

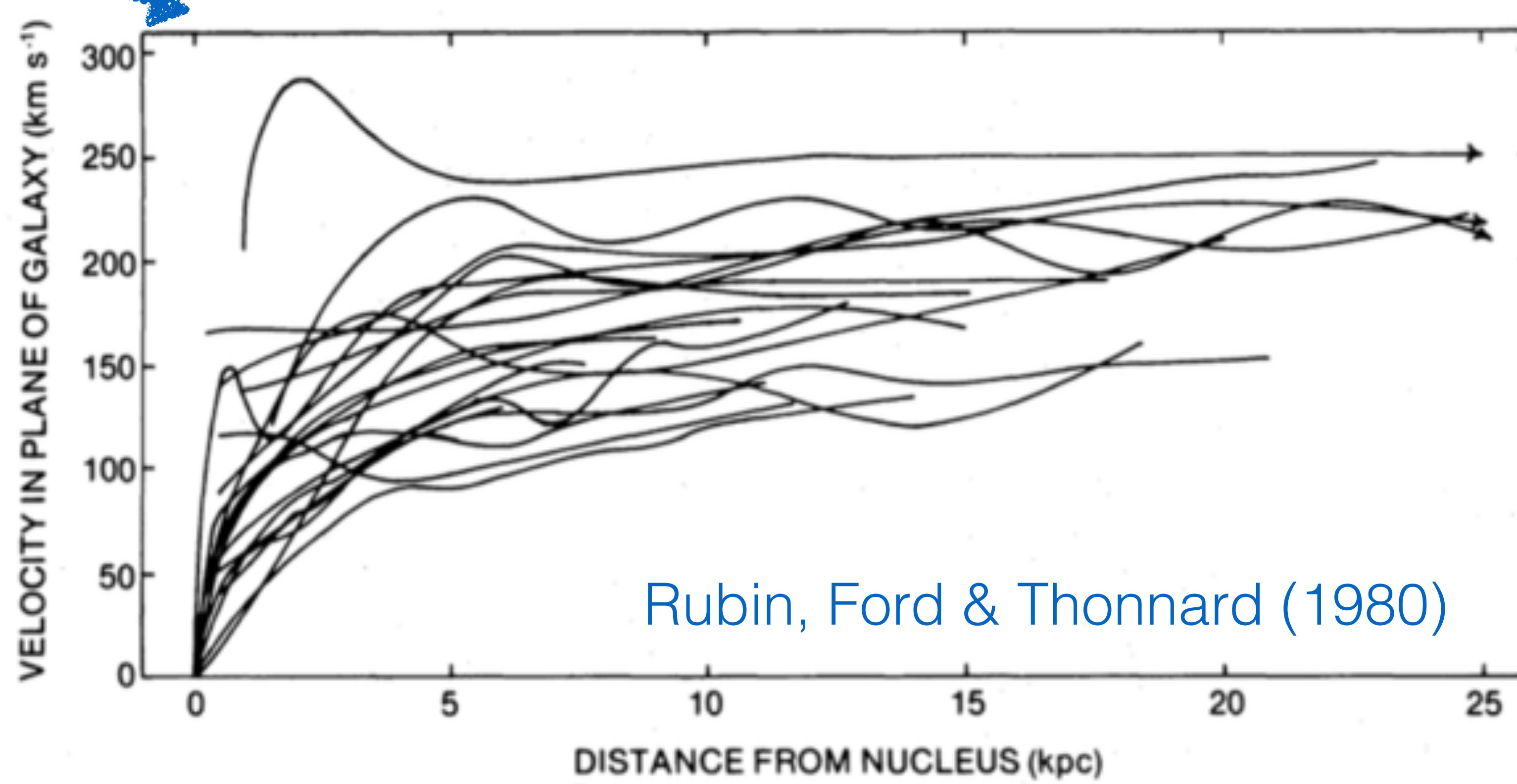
Future work

Dark Matter





Hradecky et al. [astro-ph/0006397]



Dark Matter at the Sun's Radius

Global

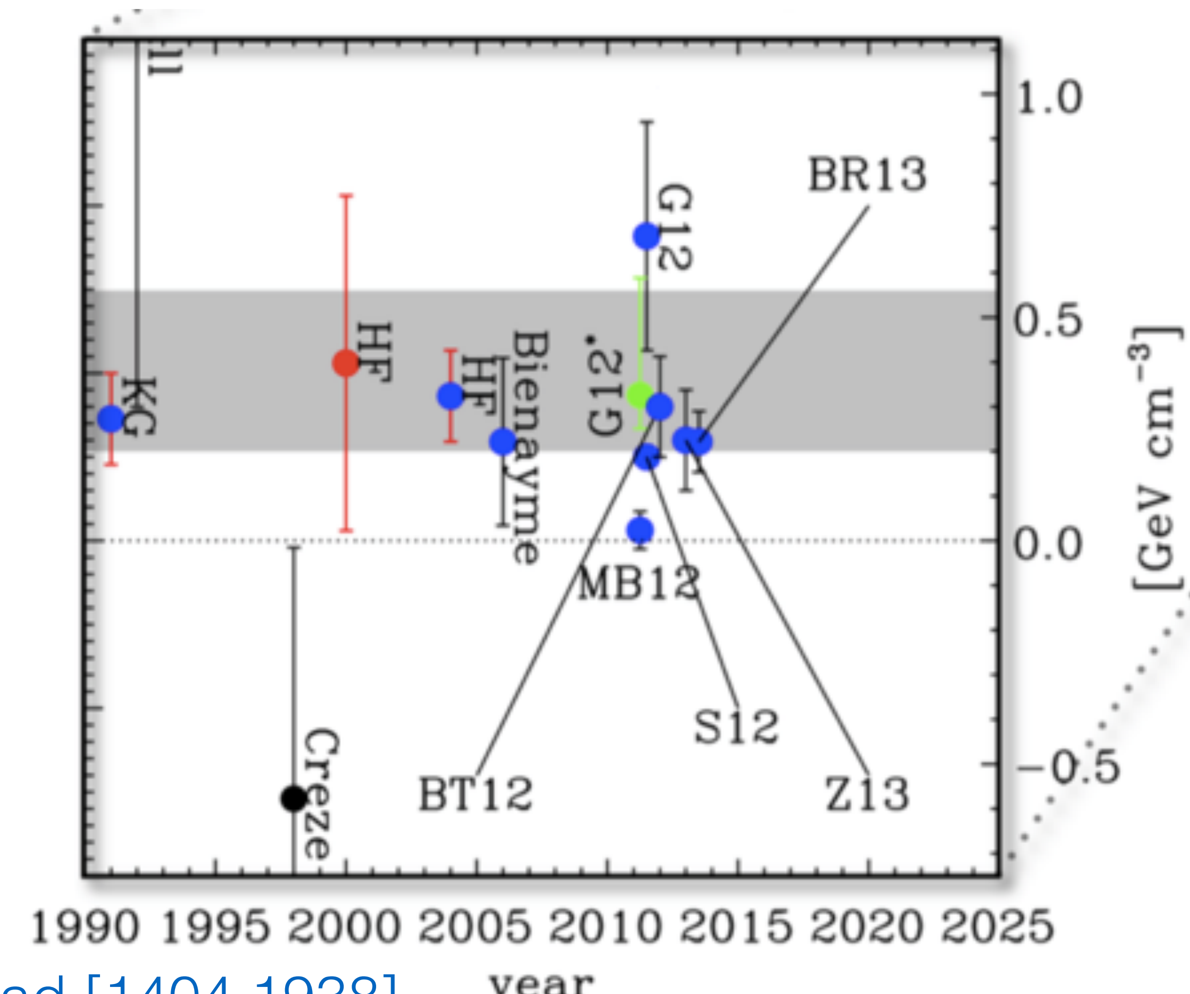
Model total mass distribution in Milky Way and extract DM density at Solar Radius (~8 kpc)

E.g. Iocco et al. [1502.03821]

Local

Estimate local DM density from kinematics of local stars (assuming local disk equilibrium)

E.g. Garbari et al. [1206.0015]



Read [1404.1938]

Values in the range:

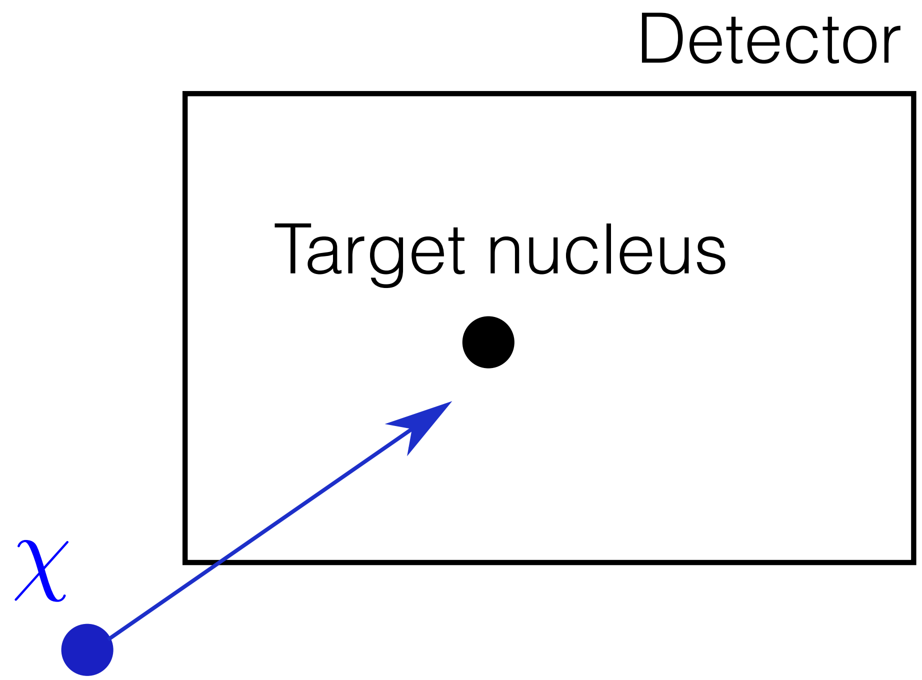
$$\rho_\chi \sim 0.2\text{–}0.8\ \mathrm{GeV\ cm^{-3}}$$

But **not** zero!

c.f. Garbari et al. [1204.3924]

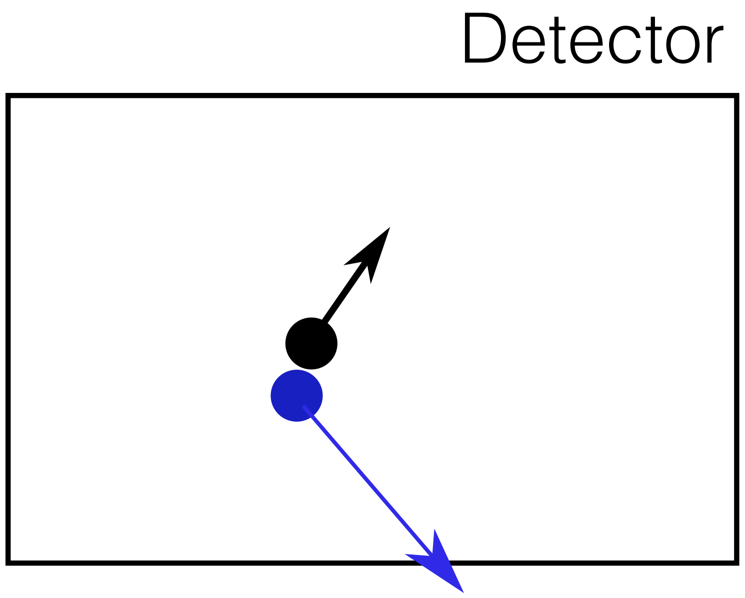
Direct detection

$$m_\chi \gtrsim 1\ \mathrm{GeV}$$
$$v \sim 10^{-3}$$



Direct detection

$m_\chi \gtrsim 1$ GeV
$v \sim 10^{-3}$



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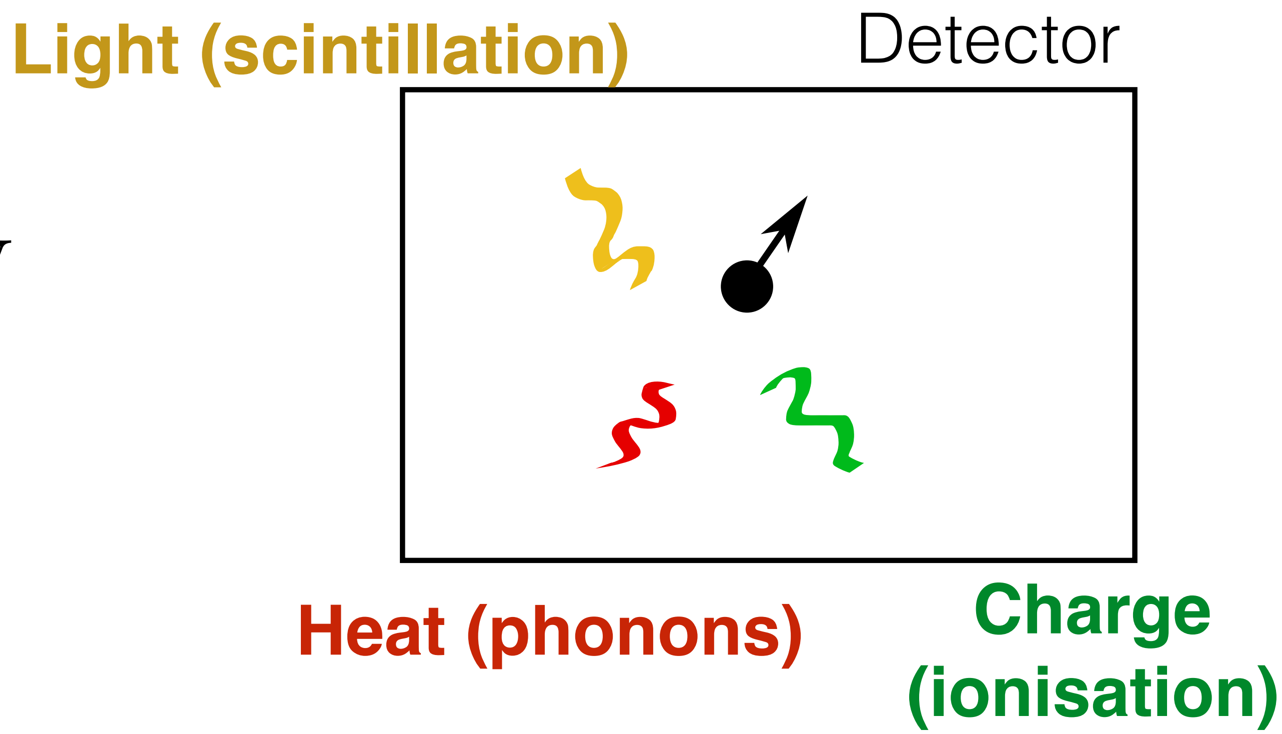
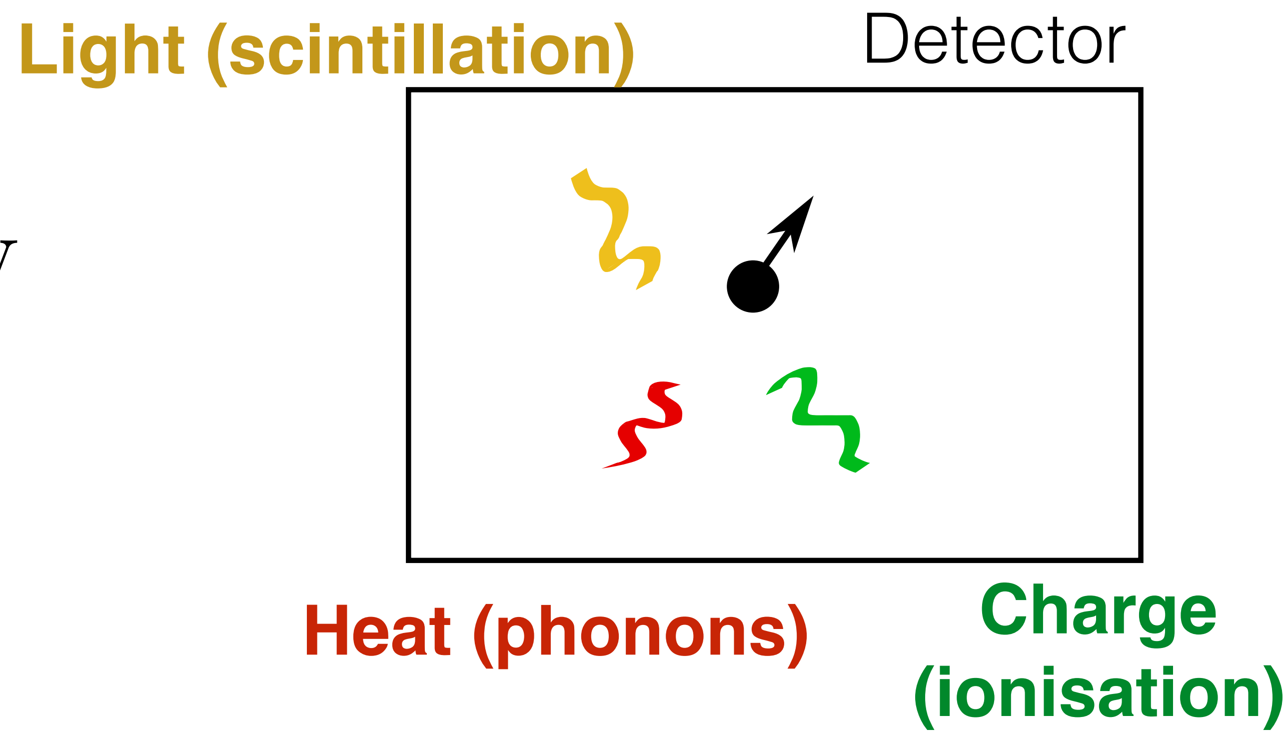


Direct detection



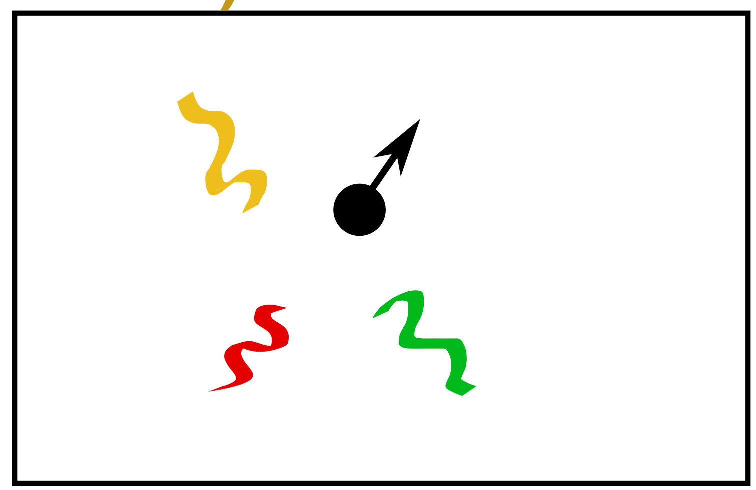
$$m_\chi \gtrsim 1\ \mathrm{GeV}$$
$$v \sim 10^{-3}$$

$$\frac{\mathrm{d}R}{\mathrm{d}E_R} = \frac{\rho_\chi}{m_\chi m_A} \int_{v_{\mathrm{min}}}^{\infty} v f(\mathbf{v}) \frac{\mathrm{d}\sigma}{\mathrm{d}E_R}\, \mathrm{d}^3\mathbf{v}$$

Include all particles with enough speed to excite recoil of energy E_R:

$$v_{\mathrm{min}} = \sqrt{\frac{m_N E_R}{2\mu_{\chi N}^2}}$$

Direct detection



$$m_\chi \gtrsim 1 \text{ GeV}$$
$$v \sim 10^{-3}$$

$$\frac{\mathrm{d}R}{\mathrm{d}E_R} = \frac{\rho_\chi}{m_\chi m_A} \int_{v_{\min}}^{\infty} v f(\mathbf{v}) \frac{\mathrm{d}\sigma}{\mathrm{d}E_R} \mathrm{d}^3\mathbf{v}$$

Astrophysics

Particle and nuclear physics

Include all particles with enough speed to excite recoil of energy E_R:

$$v_{\min} = \sqrt{\frac{m_N E_R}{2\mu_{\chi N}^2}}$$

Direct detection

$$m_\chi \gtrsim 1 \text{ GeV}$$
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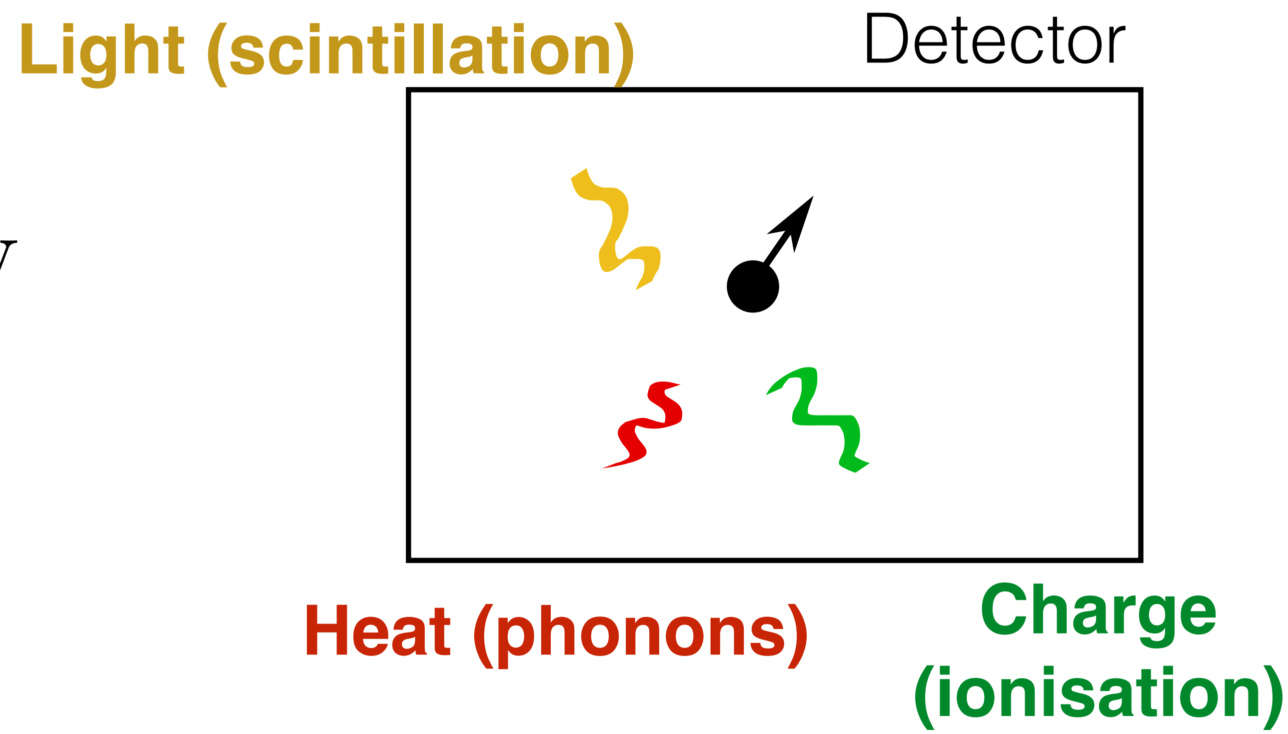


$$\frac{\mathrm{d}R}{\mathrm{d}E_R} = \frac{\rho_\chi}{m_\chi m_A} \int_{v_{\min}}^{\infty} v f(\mathbf{v}) \frac{\mathrm{d}\sigma}{\mathrm{d}E_R} \mathrm{d}^3\mathbf{v}$$

Astrophysics

Particle and nuclear physics

Include all particles with enough speed to excite recoil of energy E_R:

$$v_{\min} = \sqrt{\frac{m_N E_R}{2\mu_{\chi N}^2}}$$

But plenty of alternative ideas:
DM-electron recoils [1108.5383]
Superconducting detectors [1504.07237]
Axion DM searches [1404.1455]

Particle Physics of DM (the simple picture)

Typically assume contact interactions (heavy mediators).
In the non-relativistic limit, obtain two main contributions.
Write in terms of DM-proton cross section σ^p:

$$\frac{\mathrm{d}\sigma^A}{\mathrm{d}E_R} \propto \frac{\sigma^p}{\mu_{\chi p}^2 v^2} \mathcal{C}_A F^2(E_R)$$

Form factor accounts for loss of coherence at high energy

Enhancement factor different for:

spin-independent (SI) interactions - $\mathcal{C}_A^{\mathrm{SI}} \sim A^2$

spin-dependent (SD) interactions - $\mathcal{C}_A^{\mathrm{SD}} \sim (J+1)/J$

Interactions which are higher order in v are possible - see later…

Astrophysics of DM (the simple picture)

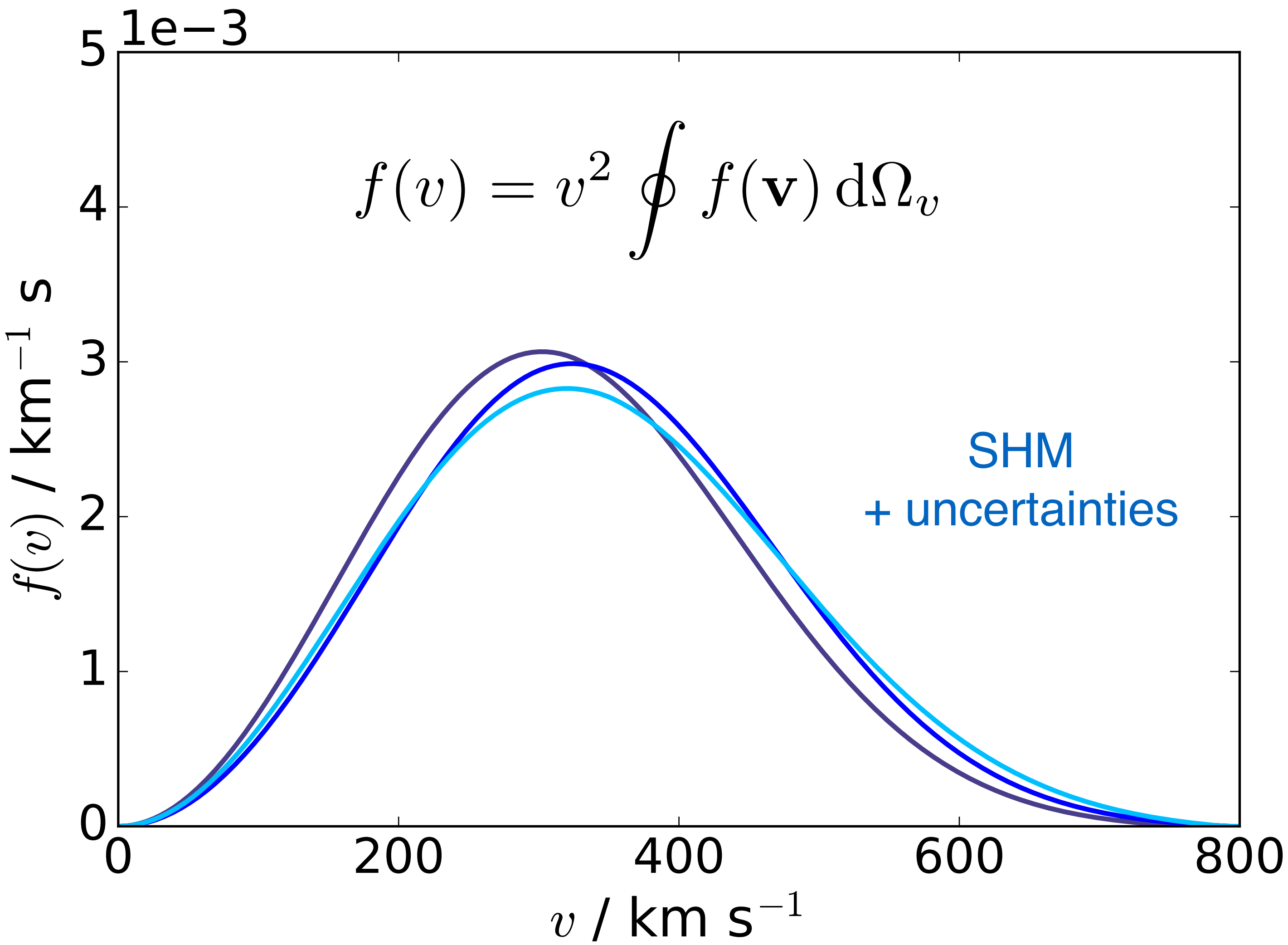
Standard Halo Model (SHM) is typically assumed: isotropic, spherically symmetric distribution of particles with $\rho(r) \propto r^{-2}$.

Leads to a Maxwell-Boltzmann (MB) distribution,

$$f_{\mathrm{Lab}}(\mathbf{v}) = (2\pi\sigma_v^2)^{-3/2} \exp\left[-\frac{(\mathbf{v}-\mathbf{v}_\mathrm{e})^2}{2\sigma_v^2}\right] \Theta(|\mathbf{v}-\mathbf{v}_\mathrm{e}| - v_\mathrm{esc})$$

which is well matched in some hydro simulations.

[1601.04707, 1601.04725, 1601.05402]



$\mathbf{v}_e$ - Earth's Velocity

$v_e \sim 220 - 250\,\mathrm{km\,s^{-1}}$

$\sigma_v \sim 155 - 175\,\mathrm{km\,s^{-1}}$

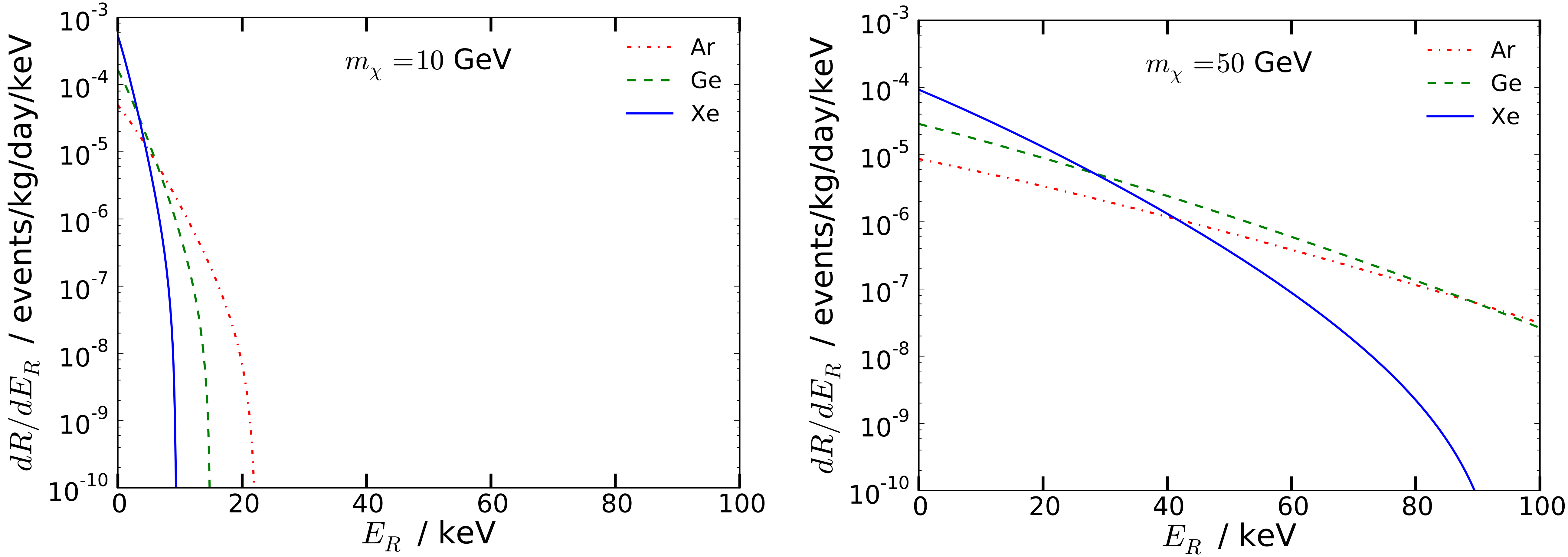
Feast et al. [astro-ph/9706293], Bovy et al. [1209.0759]

$v_\mathrm{esc} = 533^{+54}_{-41}\,\mathrm{km\,s^{-1}}$

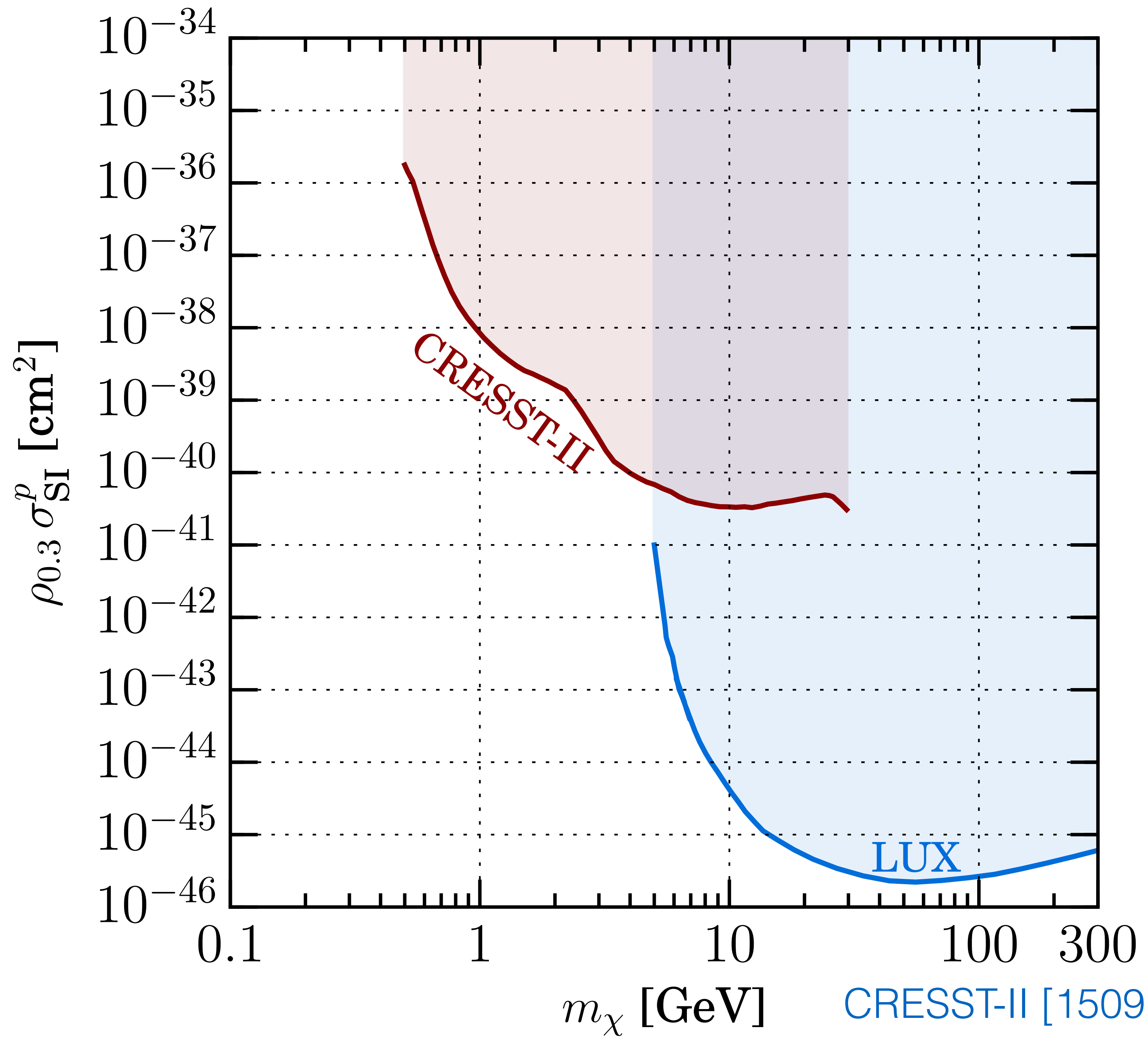
Piffl et al. (RAVE) [1309.4293]

The final event rate

SI interactions, SHM distribution



The current landscape



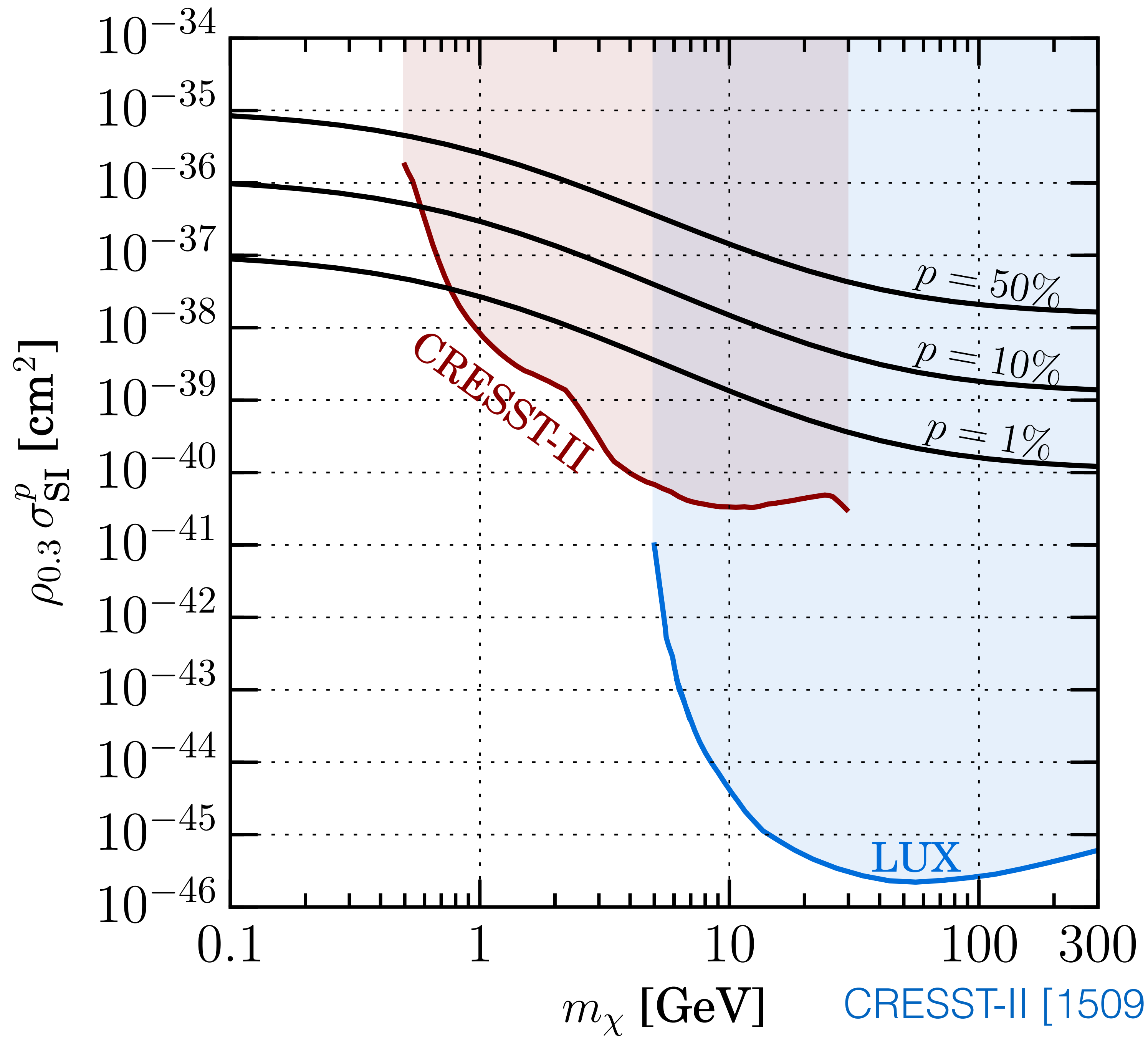
CRESST-II [1509.01515]

LUX [1608.07648]

+ many others…

How big is the probability of scattering in the Earth?

The current landscape



CRESST-II [1509.01515]

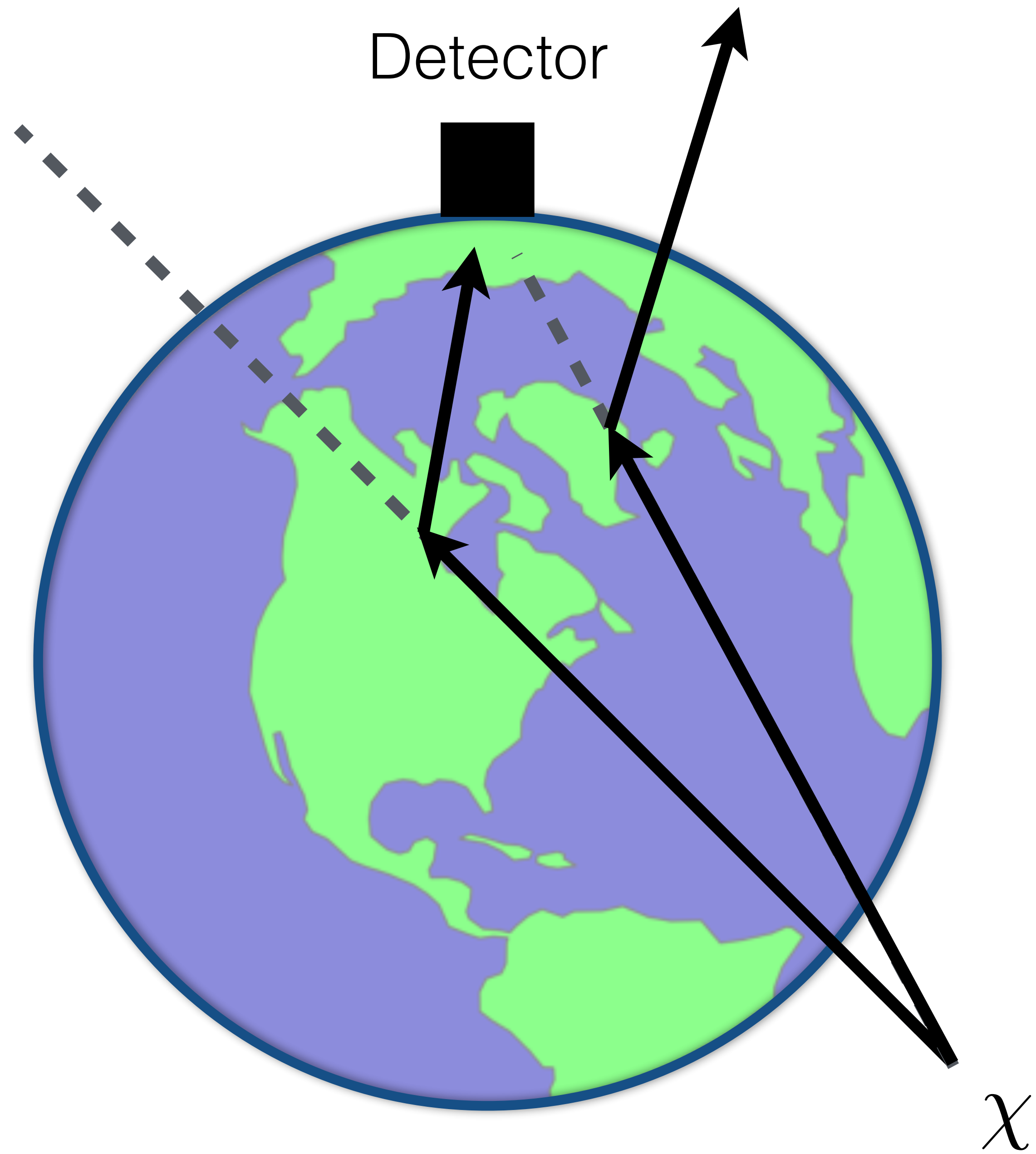
LUX [1608.07648]

+ many others…

What effect can DM scattering in the Earth have?

Earth-Shadowing

Earth-Scattering Calculation



Assuming DM mean free path

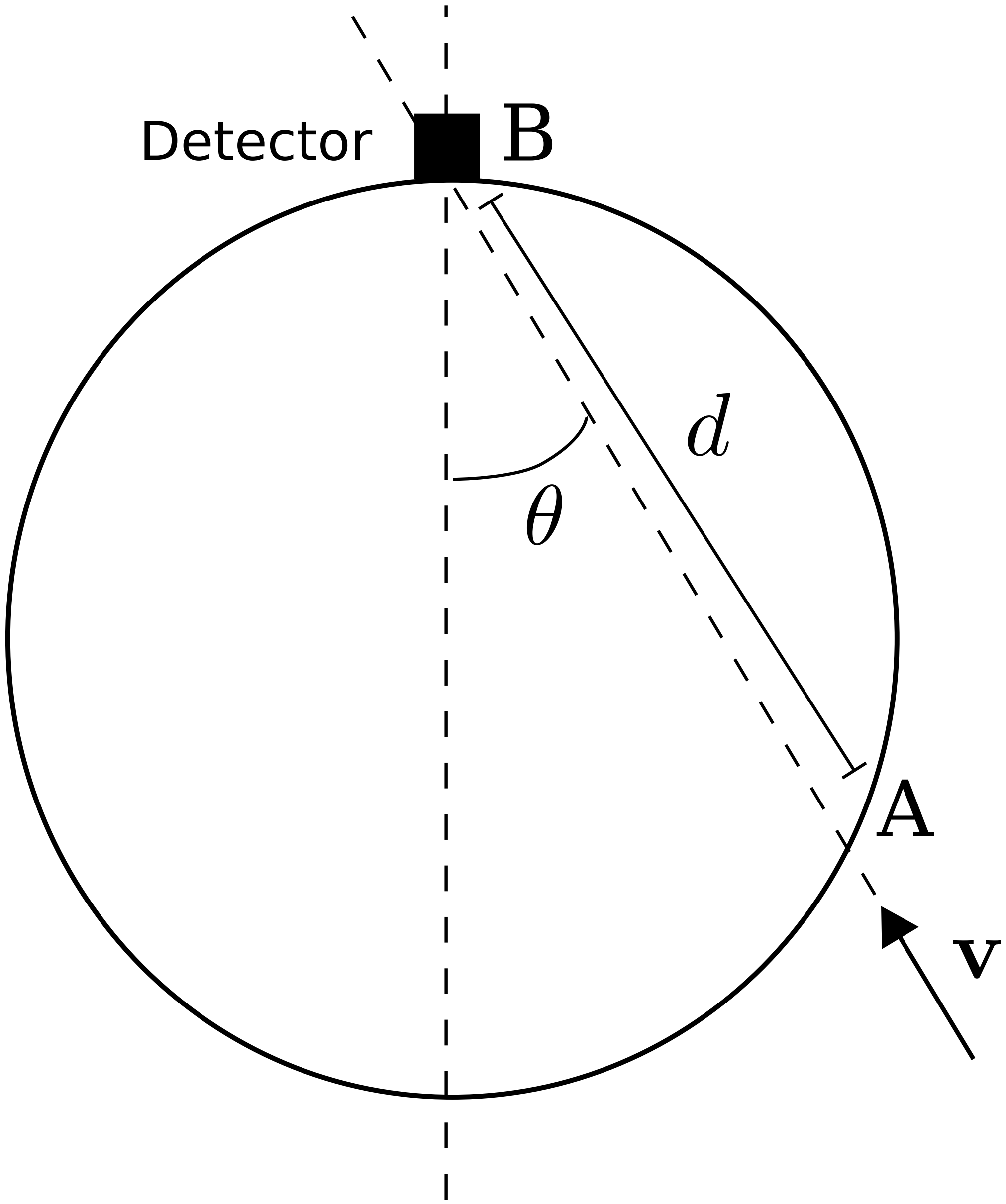
$$\lambda \gtrsim R_E$$

Total DM velocity distribution: $f(\mathbf{v}) = f_0(\mathbf{v}) - f_A(\mathbf{v}) + f_D(\mathbf{v})$

Attenuation

$$\mathbf{v} = (v, \cos\theta, \phi)$$

$$\lambda(v)^{-1} = n\,\sigma(v)$$

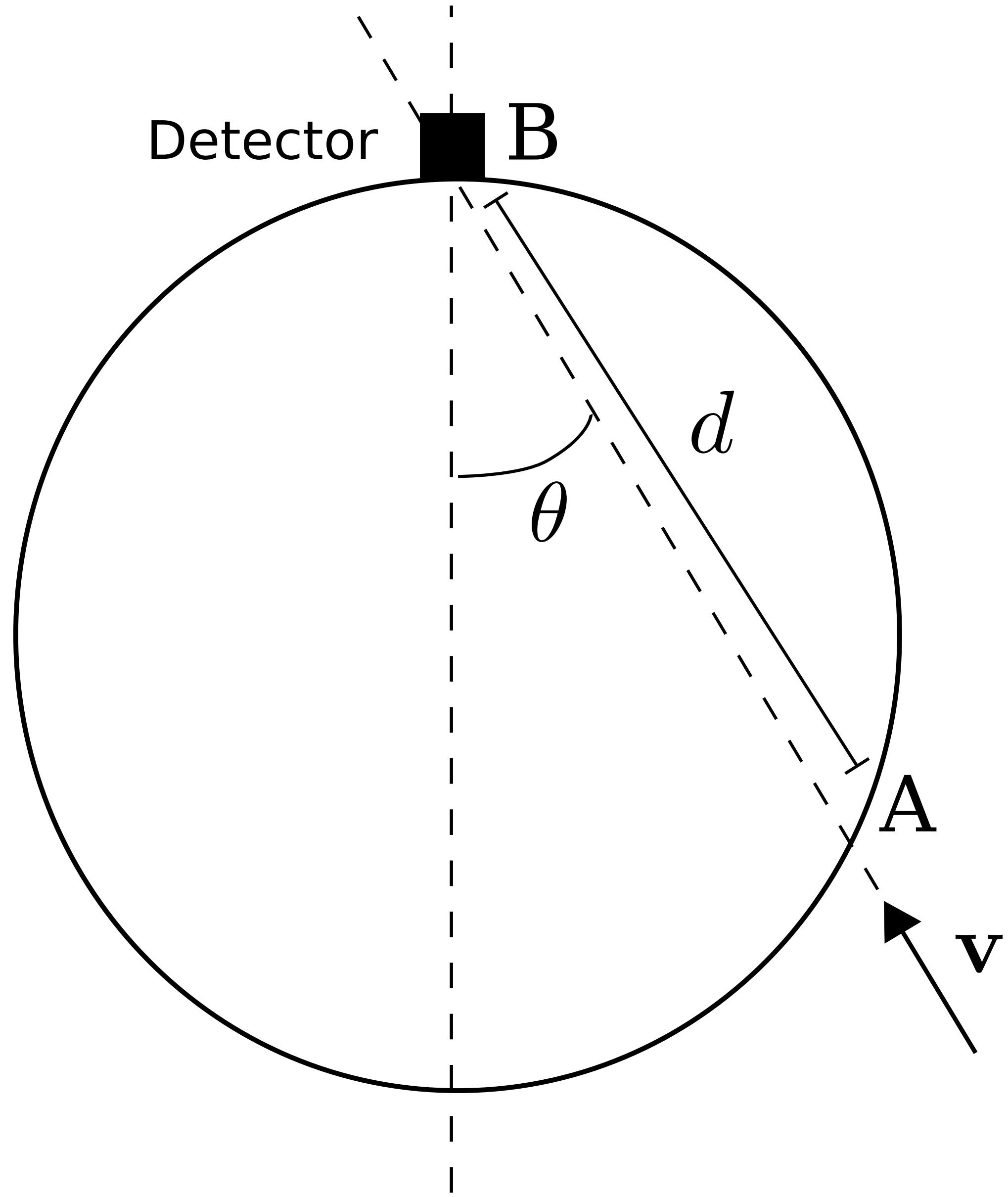


$$f_0(\mathbf{v}) - f_A(\mathbf{v}) = f_0(\mathbf{v}) \exp\left[-\frac{d(\cos\theta)}{\lambda(v)}\right]$$

Attenuation

$\mathbf{v} = (v, \cos\theta, \phi)$

$\bar{\lambda}(v)^{-1} = \bar{n}\,\sigma(v)$



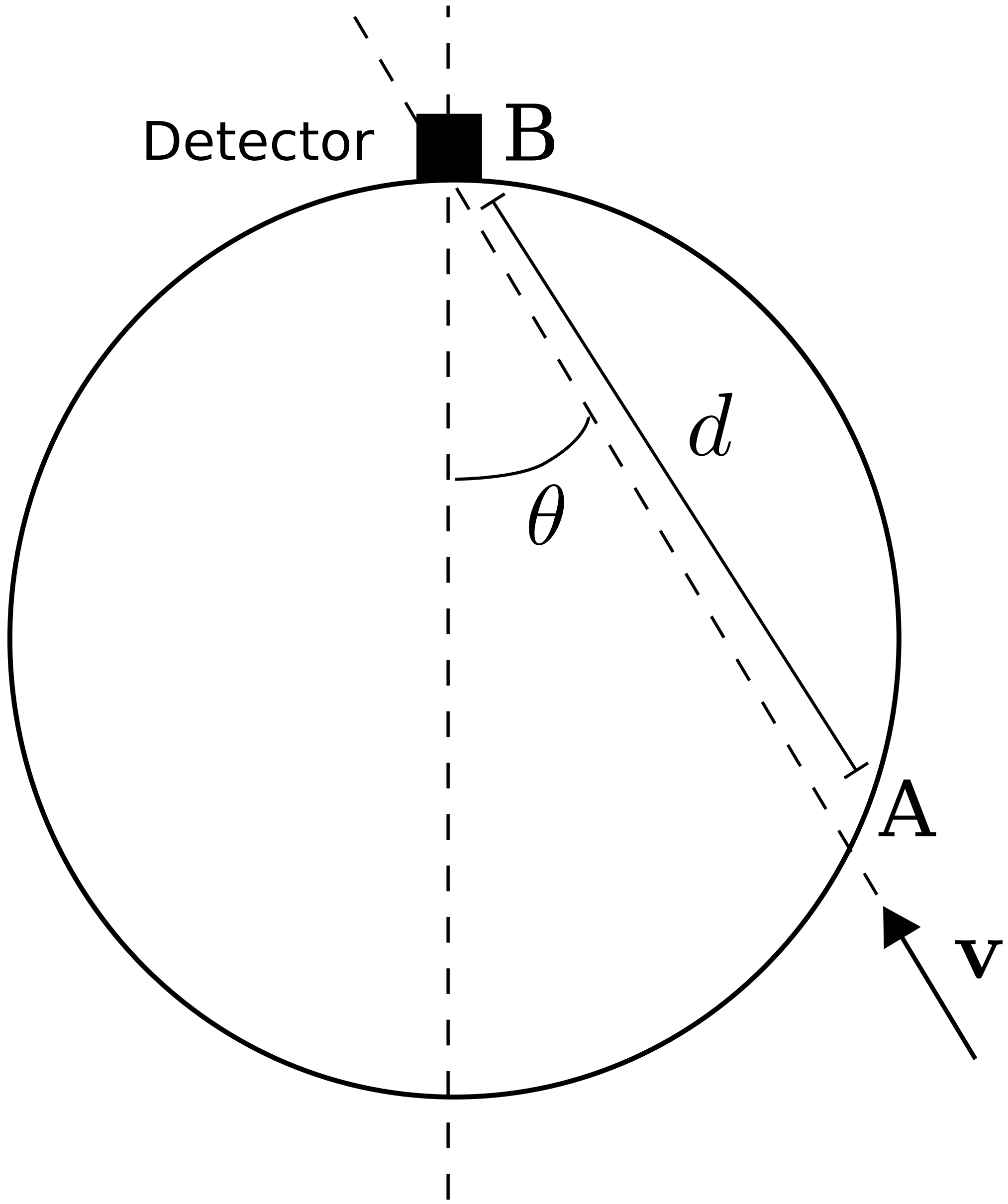
$$d_{\rm eff} = \frac{1}{\bar{n}} \int_{\rm AB} n(\mathbf{r}) \mathrm{d}l$$

$$f_0(\mathbf{v}) - f_A(\mathbf{v}) = f_0(\mathbf{v}) \exp\left[-\frac{d_{\rm eff}(\cos\theta)}{\bar{\lambda}(v)}\right]$$

Attenuation

$\mathbf{v} = (v, \cos\theta, \phi)$

$\bar{\lambda}_i(v)^{-1} = \bar{n}_i\, \sigma(v)$



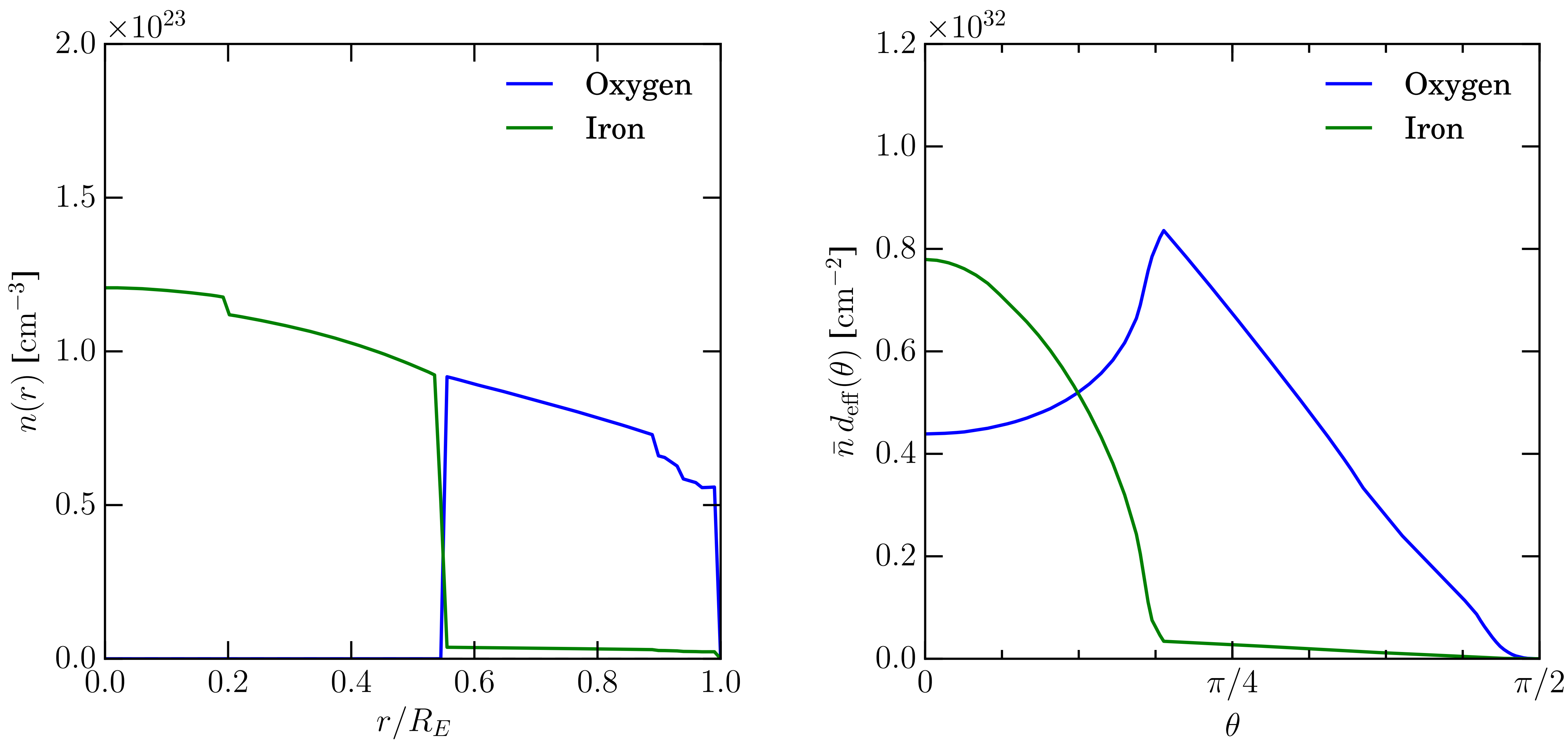
$$d_{\mathrm{eff},i} = \frac{1}{\bar{n}_i} \int_{\mathrm{AB}} n_i(\mathbf{r}) \mathrm{d}l$$

$$f_0(\mathbf{v}) - f_A(\mathbf{v}) = f_0(\mathbf{v}) \exp\left[- \sum_i^{\mathrm{species}} \frac{d_{\mathrm{eff},i}(\cos\theta)}{\bar{\lambda}_i(v)} \right]$$

Sum over 8 most abundant elements in the Earth: O, Si, Mg, Fe, Ca, Na, S, Al

Effective Earth-crossing distance

Most scattering comes from Oxygen (in the mantle) and Iron (in the core)

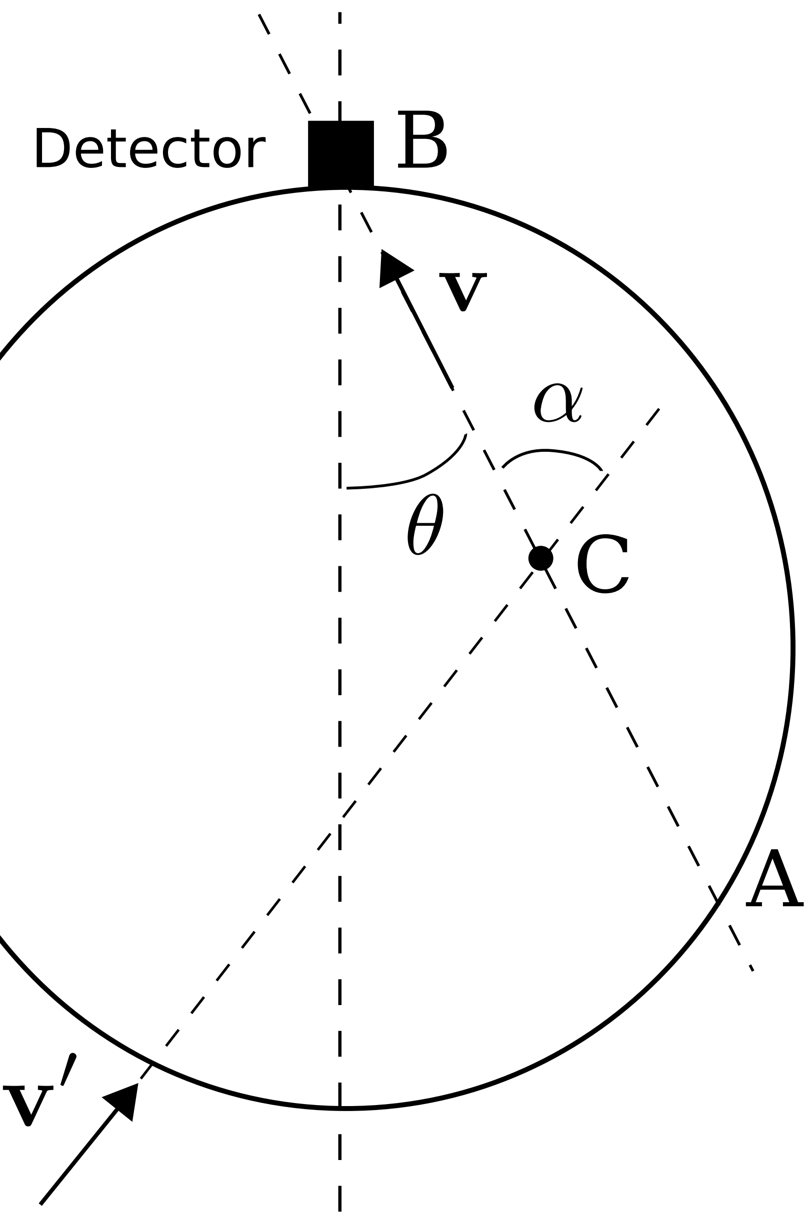


NB: little Earth-scattering for spin-dependent interactions

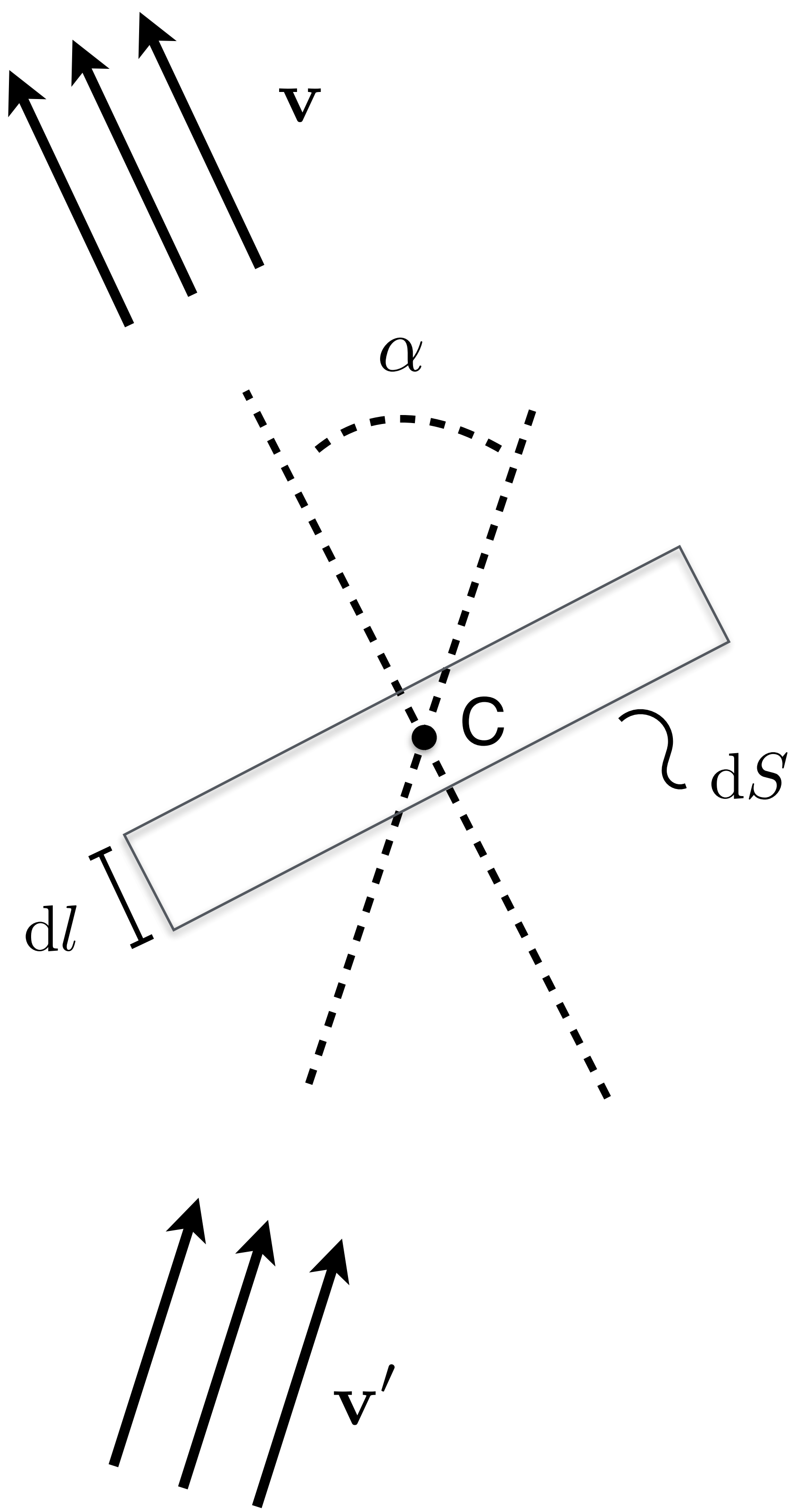
Deflection

$$\mathbf{v}' = (v', \cos\theta', \phi')$$
$$\mathbf{v} = (v, \cos\theta, \phi)$$

$$\bar{\lambda}_i(v)^{-1} = \bar{n}_i\,\sigma(v)$$



Deflection



Rate of particles entering the region:

$$n_\chi \, f_0(\mathbf{v}') \, v' \cos\alpha \, \mathrm{d}S \, \mathrm{d}^3\mathbf{v}'$$

Probability of scattering in the region:

$$\frac{\mathrm{d}l}{\lambda_i(\mathbf{r}, v') \cos\alpha} \, P(\mathbf{v}' \to \mathbf{v}) \, \mathrm{d}^3\mathbf{v}$$

Rate of particles leaving the region:

$$n_\chi f_D(\mathbf{v}) \, v \, \mathrm{d}S \, \mathrm{d}^3\mathbf{v}$$

↓

Deflected velocity distribution:

$$f_D(\mathbf{v}) = \frac{\mathrm{d}l}{\lambda_i(\mathbf{r}, v')} \frac{v'}{v} f_0(\mathbf{v}') P(\mathbf{v}' \to \mathbf{v}) \, \mathrm{d}^3\mathbf{v}'$$

Deflection

Deflected velocity distribution (from a single point):

$$f_D(\mathbf{v}) = \frac{\mathrm{d}l}{\lambda_i(\mathbf{r}, v')} \frac{v'}{v} f_0(\mathbf{v}') P(\mathbf{v}' \to \mathbf{v}) \, \mathrm{d}^3\mathbf{v}'$$

Probability of scattering from one velocity to another can be written:

$$\begin{aligned} P(\mathbf{v}' \to \mathbf{v}) &= \frac{1}{2\pi} \frac{1}{v^2} \delta(v - v'/\kappa_i) \, P(\cos\alpha) \\ &= \frac{1}{2\pi} \frac{v'}{v^3} \delta(v' - \kappa_i v) \, P(\cos\alpha) \end{aligned}$$

$$v'/v \equiv \kappa_i$$

fixed by kinematics
(for a given α)

Need to integrate over all incoming velocities and over all points C:

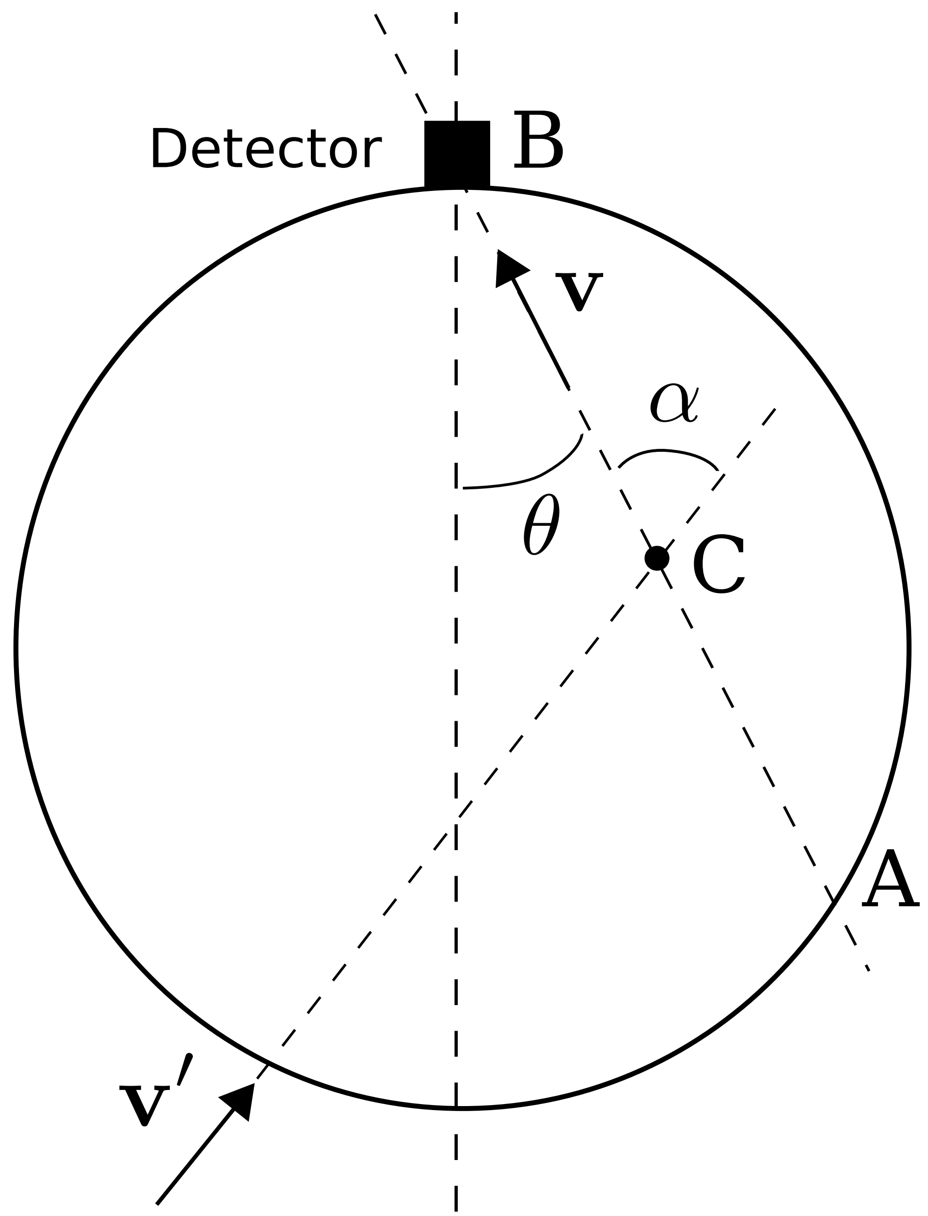
$$f_D(\mathbf{v}) = \frac{1}{2\pi} \int_{AB} \frac{\mathrm{d}l}{\lambda_i(\mathbf{r}, v')} \int \mathrm{d}^3\mathbf{v}' \, \frac{v'^2}{v^4} \delta(v' - \kappa_i v) f_0(v', \hat{\mathbf{v}}') P_i(\cos\alpha)$$

Collect everything together, and sum over Earth species…

Deflection

$$\mathbf{v}' = (v', \cos\theta', \phi')$$
$$\mathbf{v} = (v, \cos\theta, \phi)$$

$$\bar{\lambda}_i(v)^{-1} = \bar{n}_i\, \sigma(v)$$



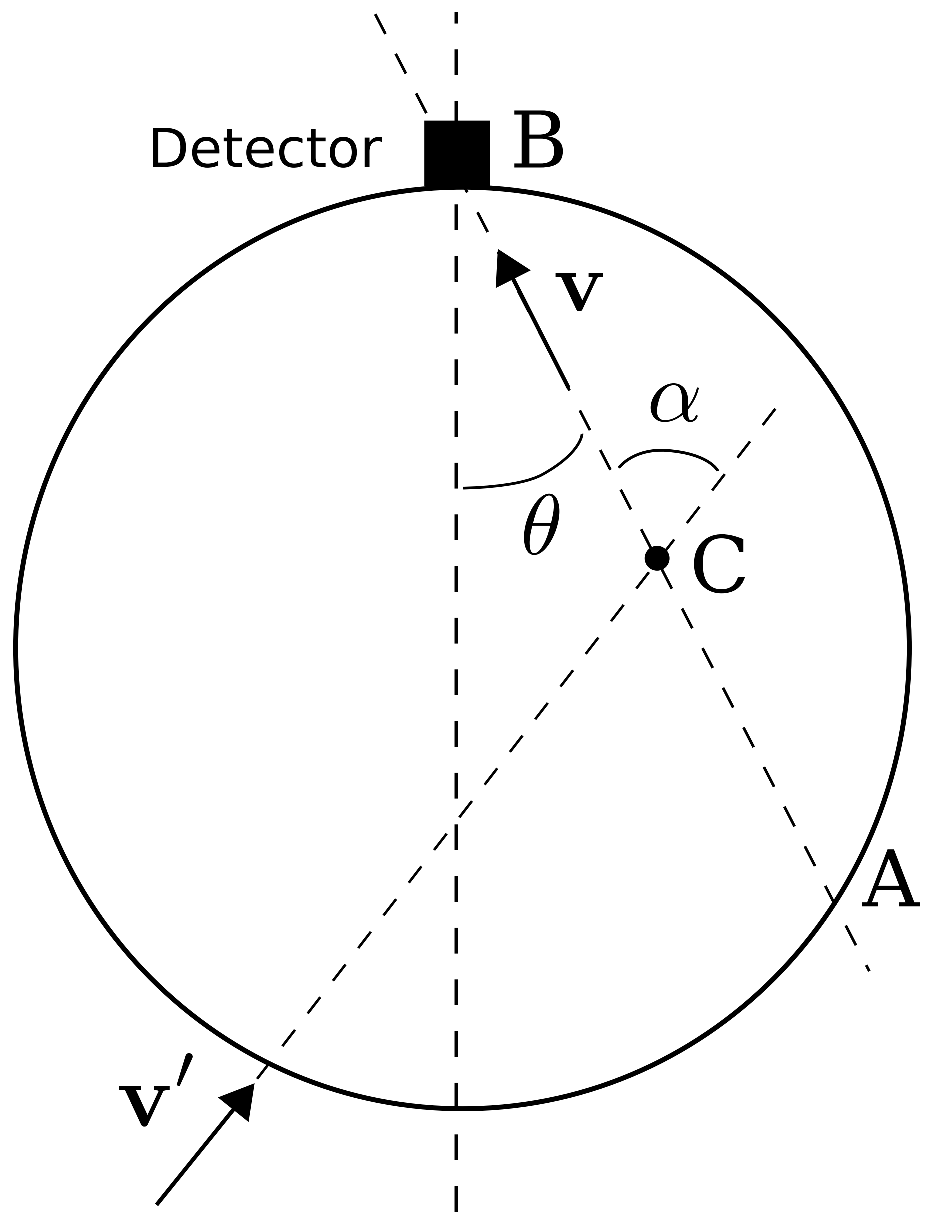
$$f_D(\mathbf{v}) = \sum_i^{\text{species}} \int \mathrm{d}^2\hat{\mathbf{v}}' \, \frac{d_{\mathrm{eff},i}(\cos\theta)}{\bar{\lambda}_i(\kappa_i v)} \frac{(\kappa_i)^4}{2\pi} f_0(\kappa_i v, \hat{\mathbf{v}}') P_i(\cos\alpha)$$

$$\kappa_i = v'/v$$

Deflection

$$\mathbf{v}' = (v', \cos\theta', \phi')$$
$$\mathbf{v} = (v, \cos\theta, \phi)$$

$$\bar{\lambda}_i(v)^{-1} = \bar{n}_i\,\sigma(v)$$



Depends on differential cross section

$$f_D(\mathbf{v}) = \sum_i^{\text{species}} \int \mathrm{d}^2\hat{\mathbf{v}}' \,\frac{d_{\text{eff},i}(\cos\theta)}{\bar{\lambda}_i(\kappa_i v)}\,\frac{(\kappa_i)^4}{2\pi} f_0(\kappa_i v, \hat{\mathbf{v}}') P_i(\cos\alpha)$$

Depends on total cross section

$$\kappa_i = v'/v$$

Non-standard DM operators

Non-relativistic Effective Field Theory (NREFT)

Write down all possible non-relativistic (NR) WIMP-*nucleon* operators which can mediate the *elastic* scattering.

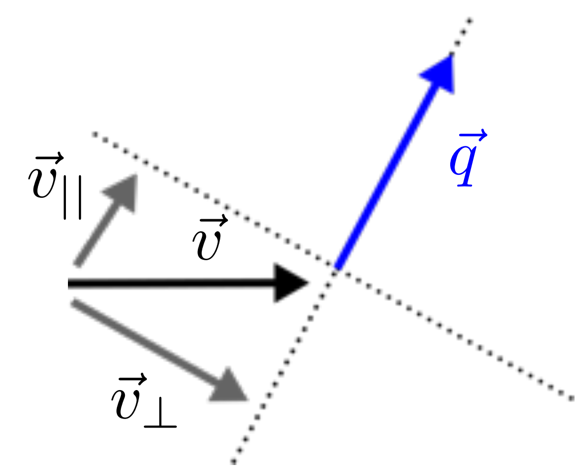
[Fan et al - 1008.1591, Fitzpatrick et al. - 1203.3542]

The building blocks of these operators are:

$$\vec{S}_\chi\,, \quad \vec{S}_N, \quad \frac{\vec{q}}{m_N}\,, \quad \vec{v}_\perp = \vec{v} + \frac{\vec{q}}{2\mu_{\chi N}}$$

The WIMP velocity operator is not Hermitian, so it can appear only through the Hermitian *transverse velocity*:

$$\vec{v}_\perp = \vec{v} + \frac{\vec{q}}{2\mu_{\chi N}} \qquad \Rightarrow \vec{v}_\perp \cdot \vec{q} = 0$$



NREFT operator basis

Write down all operators which are Hermitian, Galilean invariant and time-translation invariant:

SI → $$\mathcal{O}_1 = 1$$

SD → $$\mathcal{O}_4 = \vec{S}_\chi \cdot \vec{S}_N$$

[1008.1591, 1203.3542, 1308.6288, 1505.03117]

NREFT operator basis

Write down all operators which are Hermitian, Galilean invariant and time-translation invariant:

SI → $\mathcal{O}_1 = 1$

$\mathcal{O}_3 = i\vec{S}_N \cdot (\vec{q} \times \vec{v}^\perp)/m_N$

SD → $\mathcal{O}_4 = \vec{S}_\chi \cdot \vec{S}_N$

$\mathcal{O}_5 = i\vec{S}_\chi \cdot (\vec{q} \times \vec{v}^\perp)/m_N$

$\mathcal{O}_6 = (\vec{S}_\chi \cdot \vec{q})(\vec{S}_N \cdot \vec{q})/m_N^2$

$\mathcal{O}_7 = \vec{S}_N \cdot \vec{v}^\perp$

$\mathcal{O}_8 = \vec{S}_\chi \cdot \vec{v}^\perp$

$\mathcal{O}_9 = i\vec{S}_\chi \cdot (\vec{S}_N \times \vec{q})/m_N$

$\mathcal{O}_{10} = i\vec{S}_N \cdot \vec{q}/m_N$

$\mathcal{O}_{11} = i\vec{S}_\chi \cdot \vec{q}/m_N$

$\mathcal{O}_{12} = \vec{S}_\chi \cdot (\vec{S}_N \times \vec{v}^\perp)$

$\mathcal{O}_{13} = i(\vec{S}_\chi \cdot \vec{v}^\perp)(\vec{S}_N \cdot \vec{q})/m_N$

$\mathcal{O}_{14} = i(\vec{S}_\chi \cdot \vec{q})(\vec{S}_N \cdot \vec{v}^\perp)/m_N$

$\mathcal{O}_{15} = -(\vec{S}_\chi \cdot \vec{q})((\vec{S}_N \times \vec{v}^\perp) \cdot \vec{q}/m_N^2$

$\vdots$

NB: two sets of operators, one for protons and one for neutrons…

[1008.1591, 1203.3542, 1308.6288, 1505.03117]

Example: Anapole DM

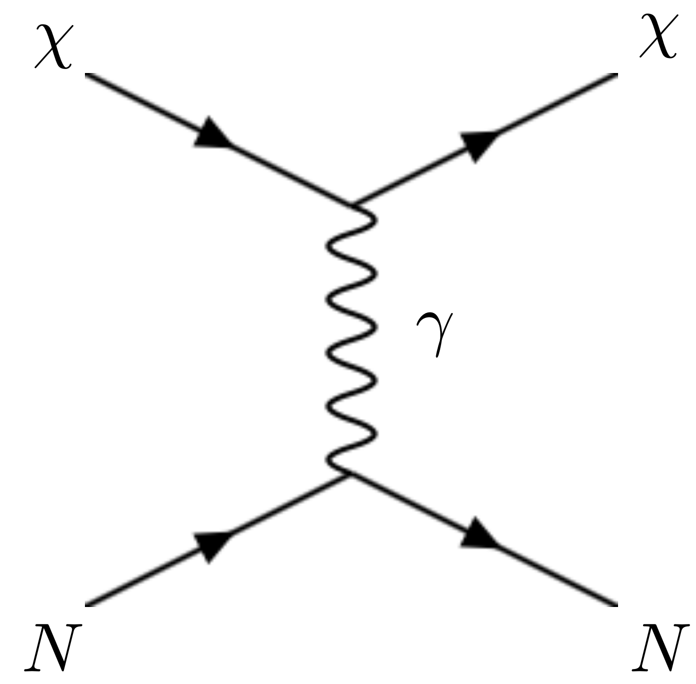
[1211.0503, 1401.4508, 1506.04454]

Lowest order interaction of Majorana DM with EM fields:

$$\mathcal{O}_A = \bar{\chi}\gamma^\mu\gamma^5\chi\,\partial^\nu F_{\mu\nu}$$

Induces an interaction with nucleons:

$$\mathcal{O}_A^{(N)} = eQ_N\bar{\chi}\gamma^\mu\gamma^5\chi\,\bar{N}\gamma_\mu N$$

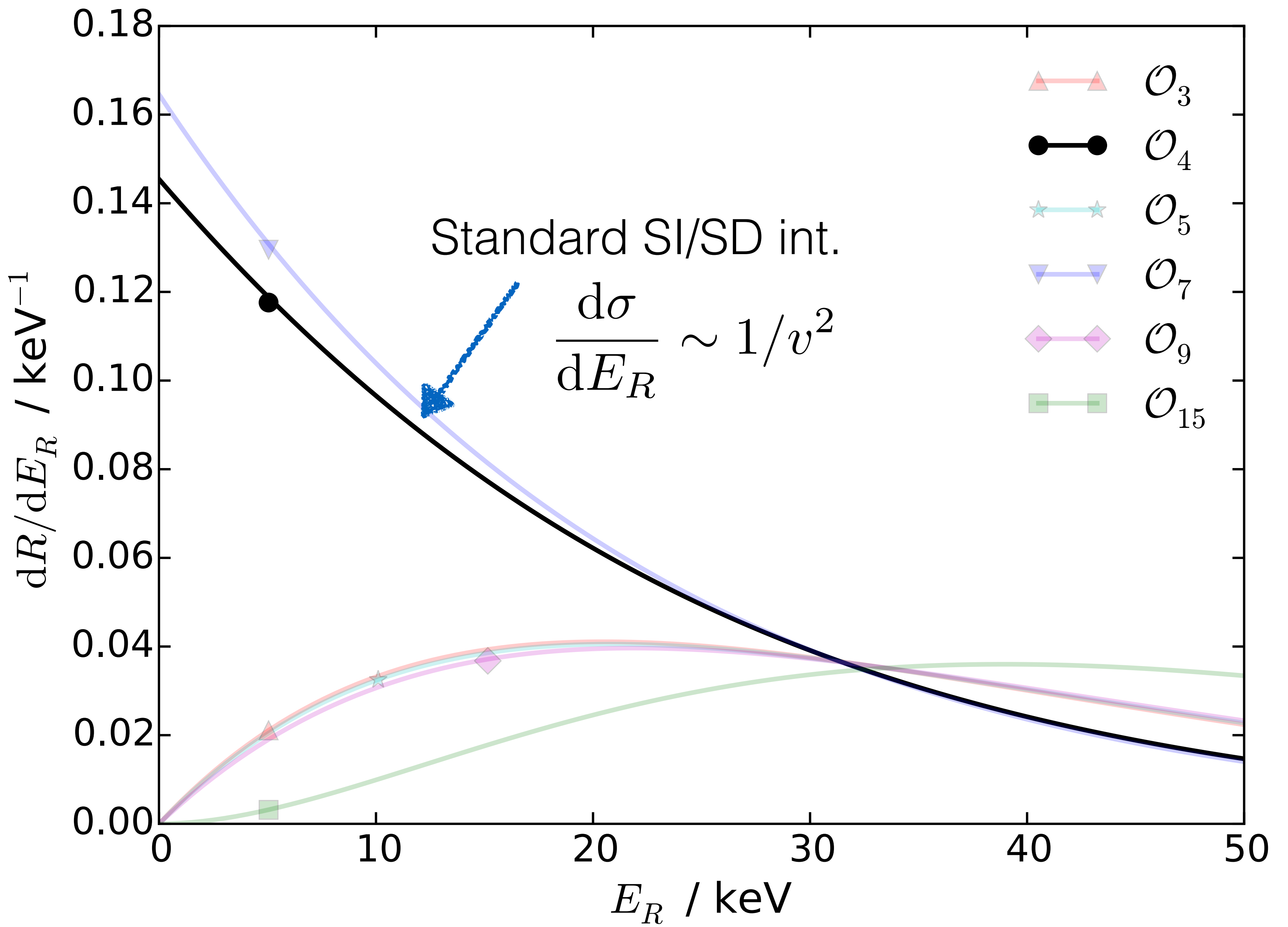


Leading to a NR matrix element:

$$\begin{aligned}\mathcal{M}_A^{(N)} &= -eQ_N m_\chi m_N \vec{S}_\chi \cdot (\vec{v}^\perp + i\vec{S}_N \times \vec{q}) \\ &= -eQ_N m_\chi m_N (\mathcal{O}_8 + \mathcal{O}_9)\end{aligned}$$

Energy spectra

$m_\chi = 100 \text{ GeV}$



Energy spectra

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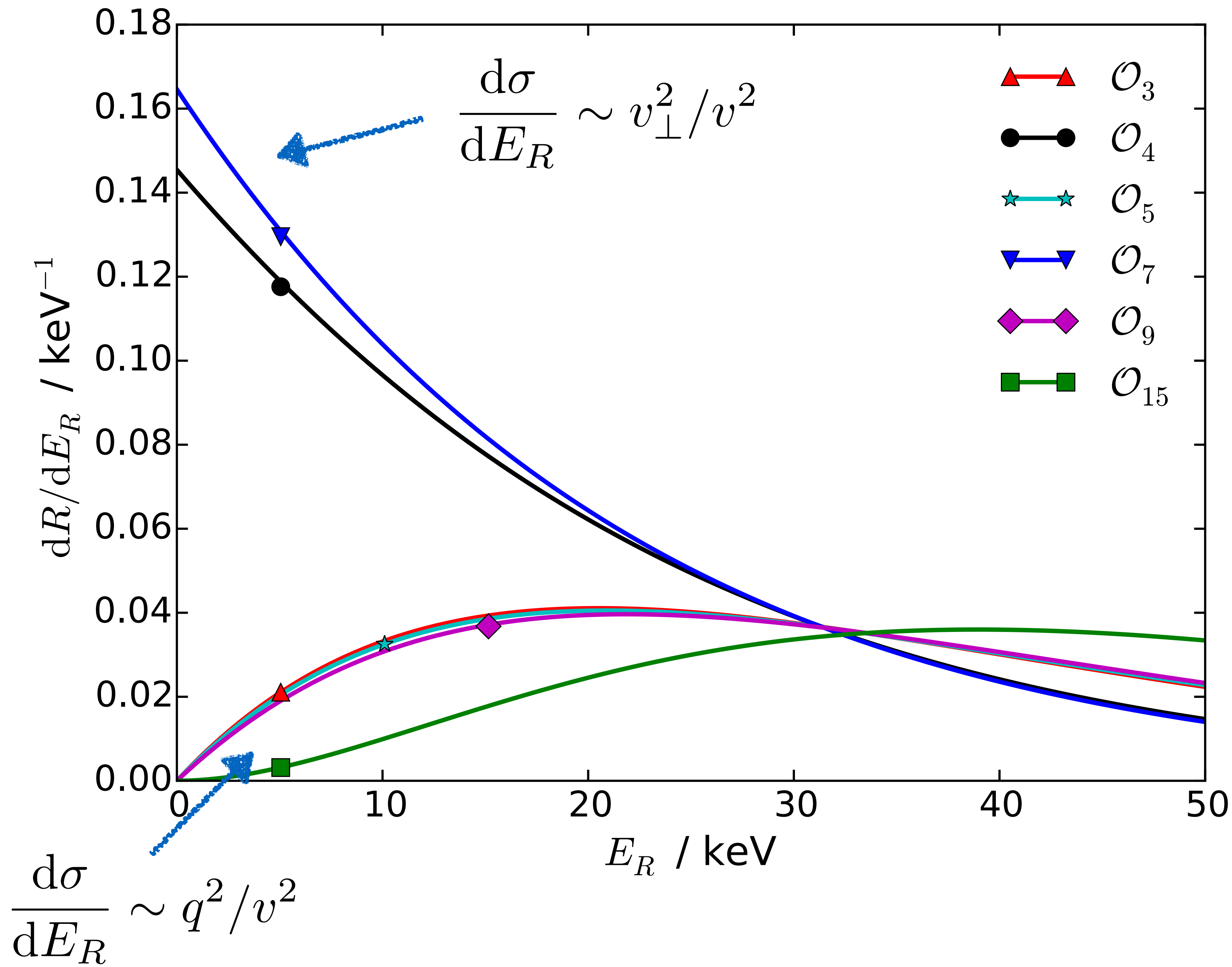
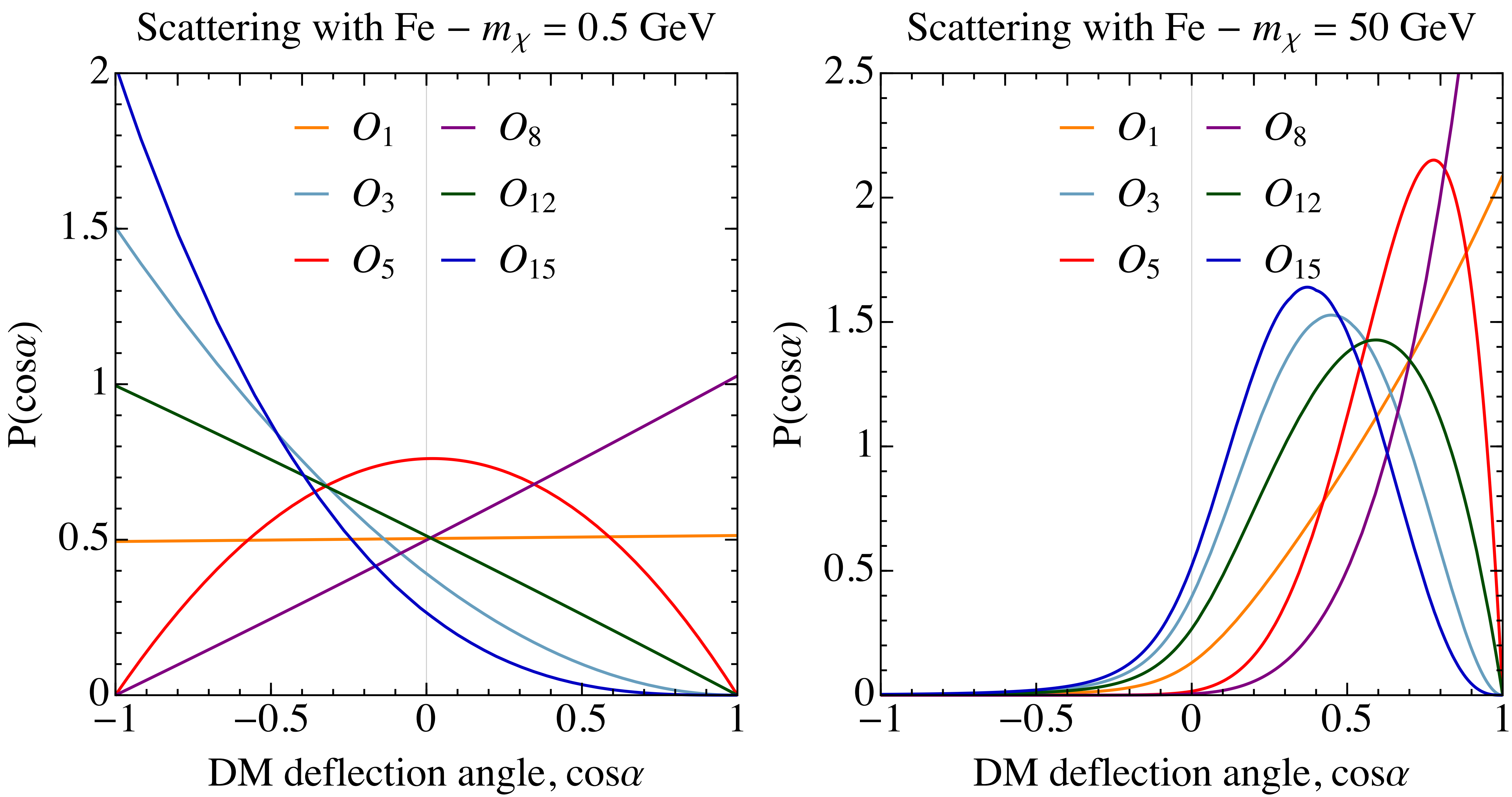


DM deflection distribution

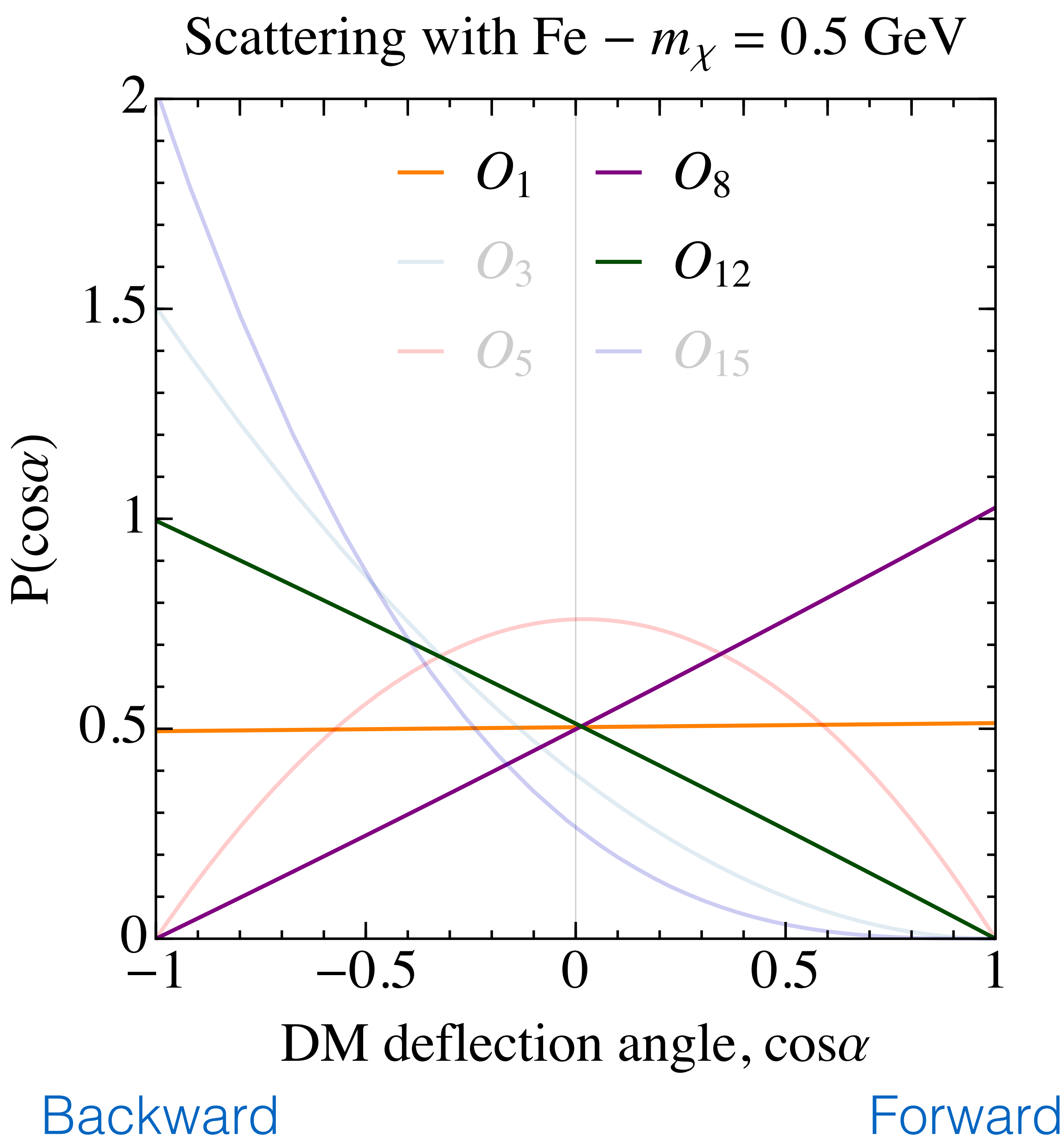
$$P(\cos\alpha) = \frac{1}{\sigma}\frac{\mathrm{d}\sigma}{\mathrm{d}E_R}\frac{\mathrm{d}E_R}{\mathrm{d}\cos\alpha}$$



Backward Forward

DM deflection distribution

$$P(\cos\alpha) = \frac{1}{\sigma}\frac{\mathrm{d}\sigma}{\mathrm{d}E_R}\frac{\mathrm{d}E_R}{\mathrm{d}\cos\alpha}$$



Standard SI

$$\mathcal{O}_1 = \mathbb{1} \Rightarrow \frac{\mathrm{d}\sigma}{\mathrm{d}E_R} \sim \frac{1}{v^2}$$

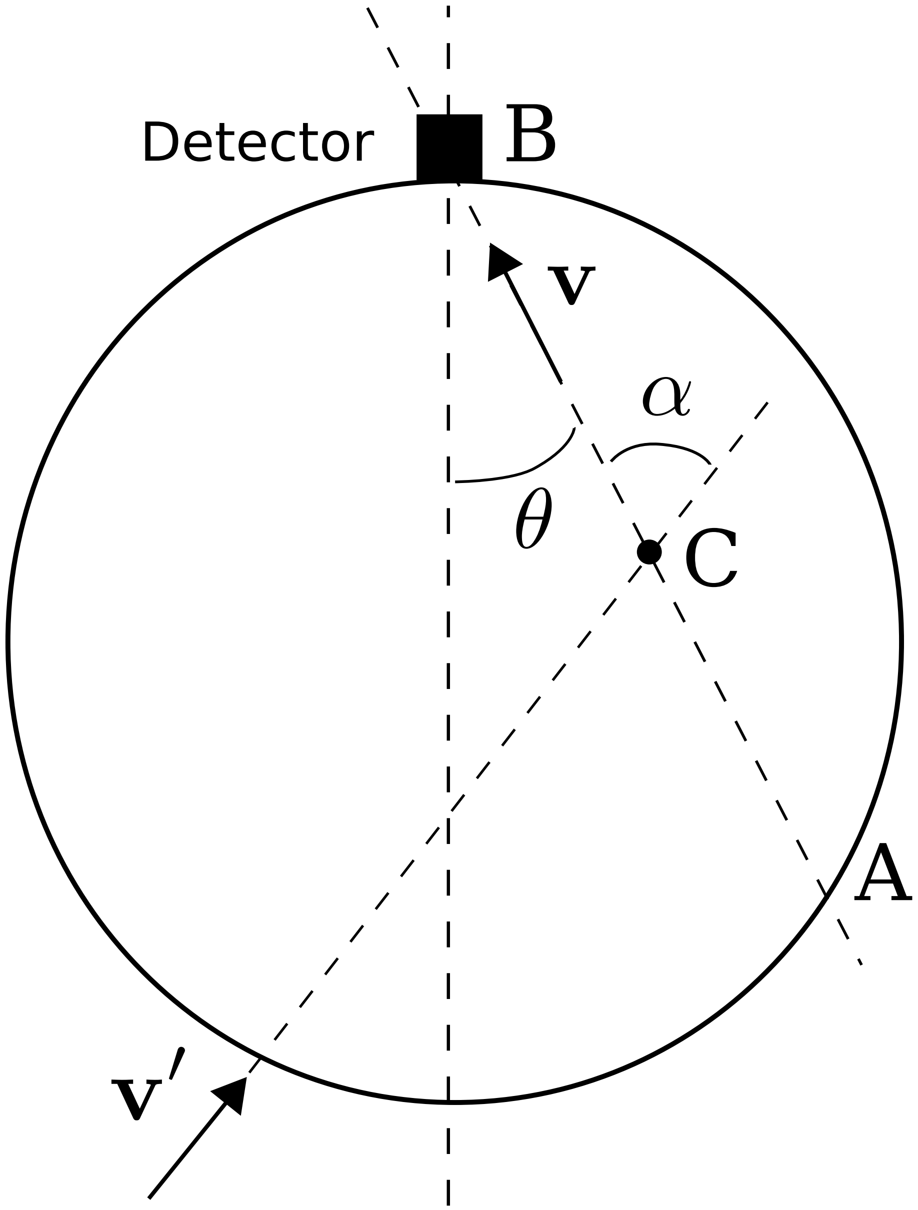
$$\mathcal{O}_8 = \vec{S}_\chi \cdot \vec{v}^\perp \Rightarrow \frac{\mathrm{d}\sigma}{\mathrm{d}E_R} \sim (1 - \frac{m_\mathcal{N} E_R}{2\mu_{\chi\mathcal{N}}^2 v^2})$$

$$\mathcal{O}_{12} = \vec{S}_\chi \cdot (\vec{S}_N \times \vec{v}^\perp) \Rightarrow \frac{\mathrm{d}\sigma}{\mathrm{d}E_R} \sim \frac{E_R}{v^2}$$

DM deflection

$$\mathbf{v}' = (v', \cos\theta', \phi')$$
$$\mathbf{v} = (v, \cos\theta, \phi)$$

$$\bar{\lambda}_i(v)^{-1} = \bar{n}_i\,\sigma(v)$$



$$f_D(\mathbf{v}) = \sum_i^{\mathrm{species}} \int \mathrm{d}^2\hat{\mathbf{v}}' \, \frac{d_{\mathrm{eff},i}(\cos\theta)}{\bar{\lambda}_i(\kappa_i v)} \frac{(\kappa_i)^4}{2\pi} f_0(\kappa_i v, \hat{\mathbf{v}}') P_i(\cos\alpha)$$

$$\kappa_i = v'/v$$

EARTHSHADOW Code

EARTHSHADOW code (will be) available online at:
github.com/bradkav/EarthShadow

Including routines, numerical results, plots and animations…

videos	Update README.md	just now
.git_ignore	Create .git_ignore	a day ago
LICENSE	Initial commit	2 months ago
README.md	Update README.md	a day ago

README.md

EarthShadow

Tables of results

In the 'results' folder, you will find a number of tabulated results. Each file is named as `Speeddist_op=[operator]_mx=0.50_gam=[X.YZ]_ps=0.10.dat`, where `[operator]` is either `1`, `8` or `12`. The results are for a DM mass of 0.5 GeV and a scattering probability of 10%. The value `[X.YZ]` is the angle gamma in units of pi (so 1.00 is gamma = pi).

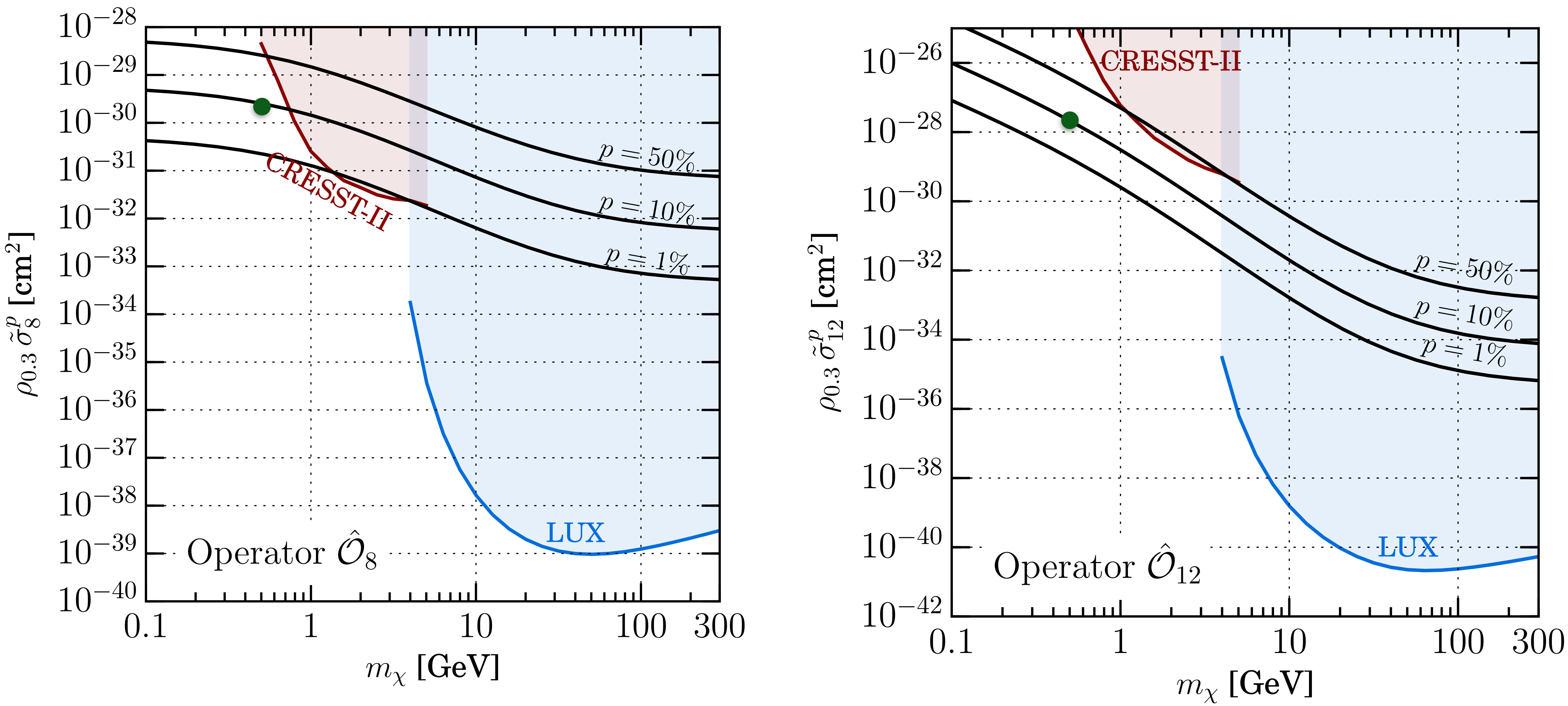
Each data file contains 4 columns:

v [km/s]	f(v) [s/km]	f_A(v) - f(v) [s/km]	f_D(v) [s/km]

Results

Constraints on NREFT operators

Focus on SI operator (O_1), as well as O_8 and O_{12}:

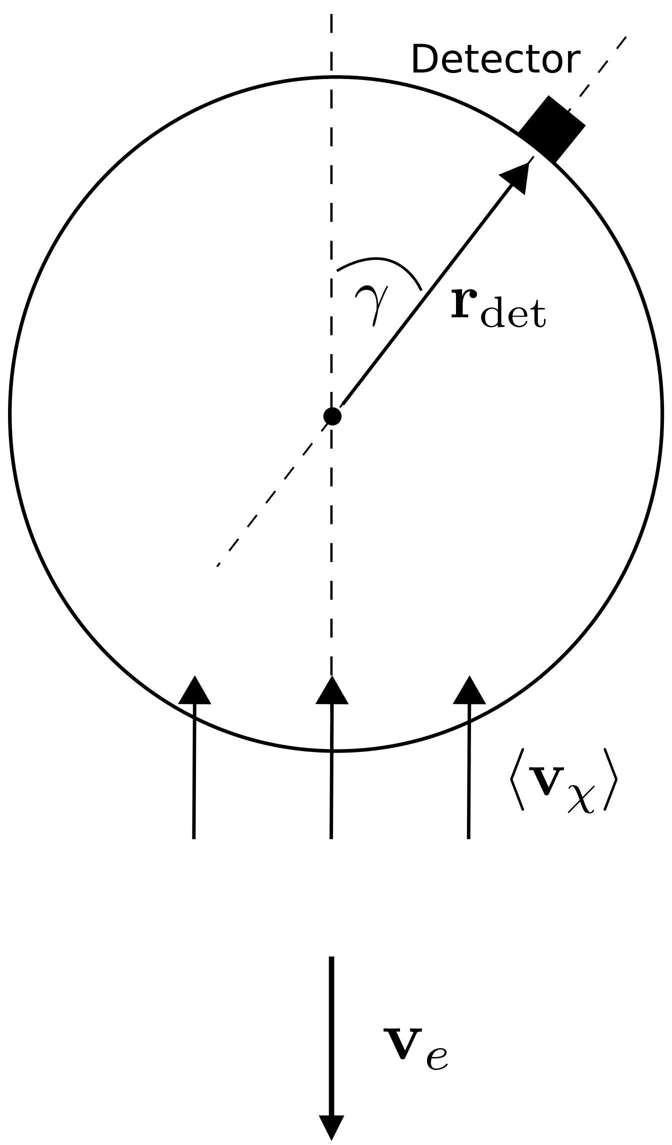


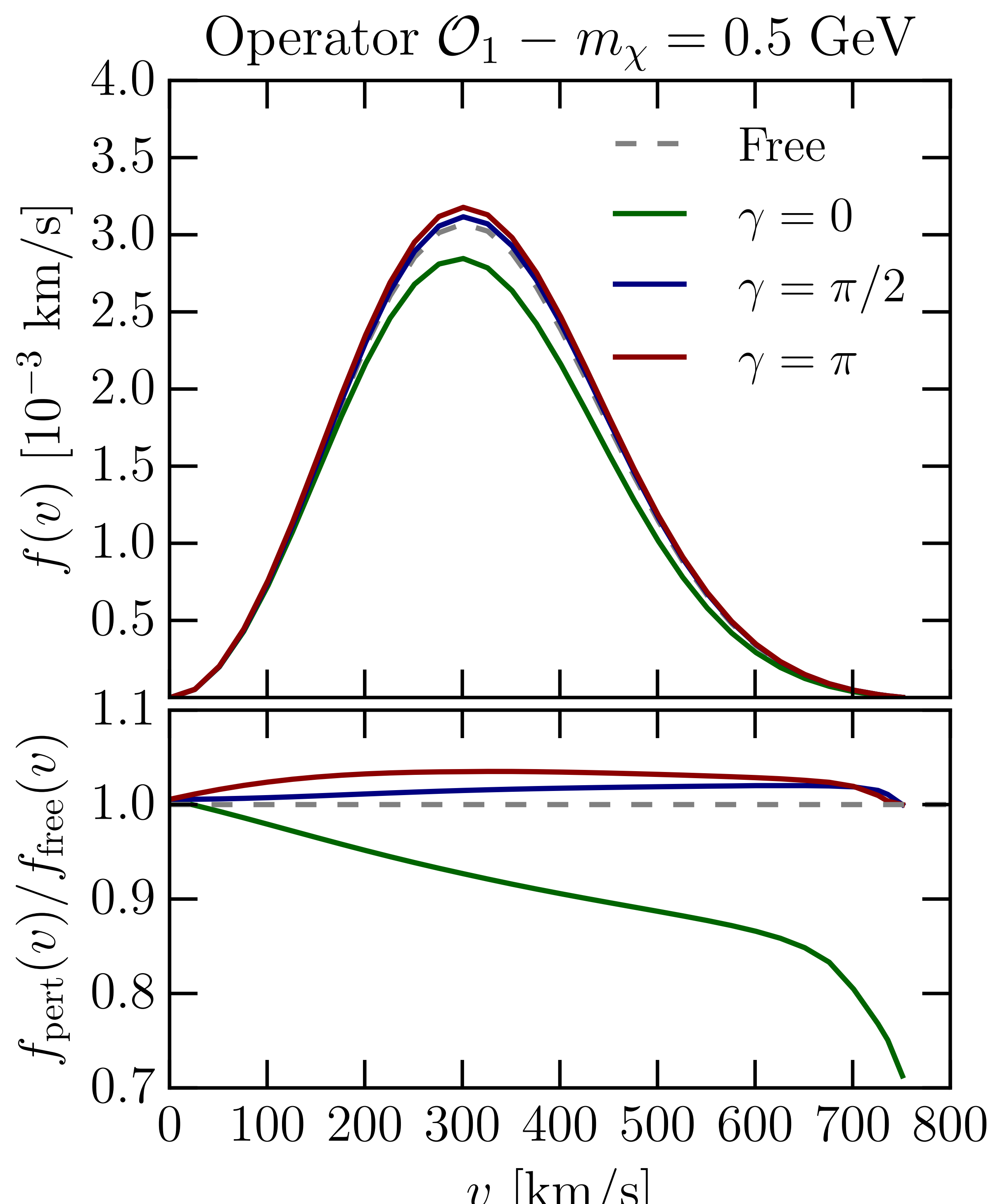
Focus on low mass DM: $m_\chi = 0.5$ GeV

Fix couplings to give 10% probability of scattering

Speed Distribution - Operator 1

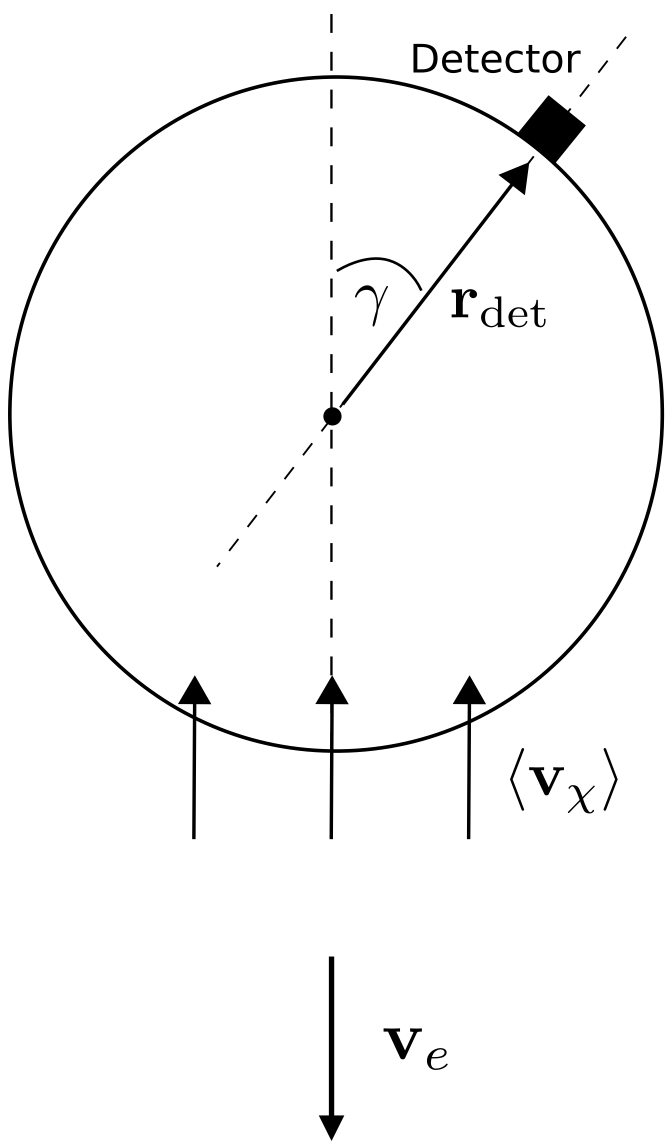
Calculate DM speed distribution after Earth scattering: $f_{\mathrm{pert}}(v)$



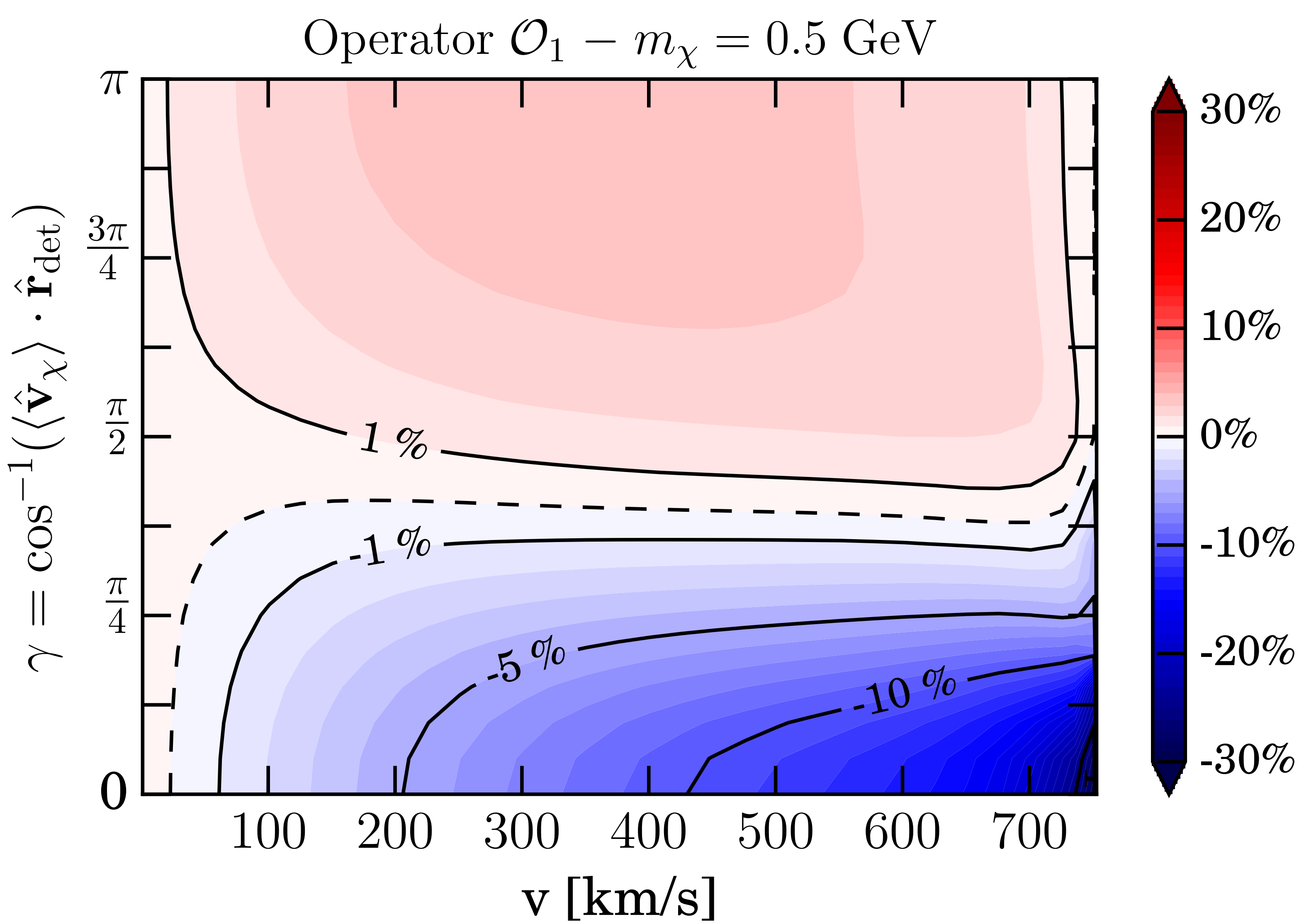


Speed Distribution - Operator 1

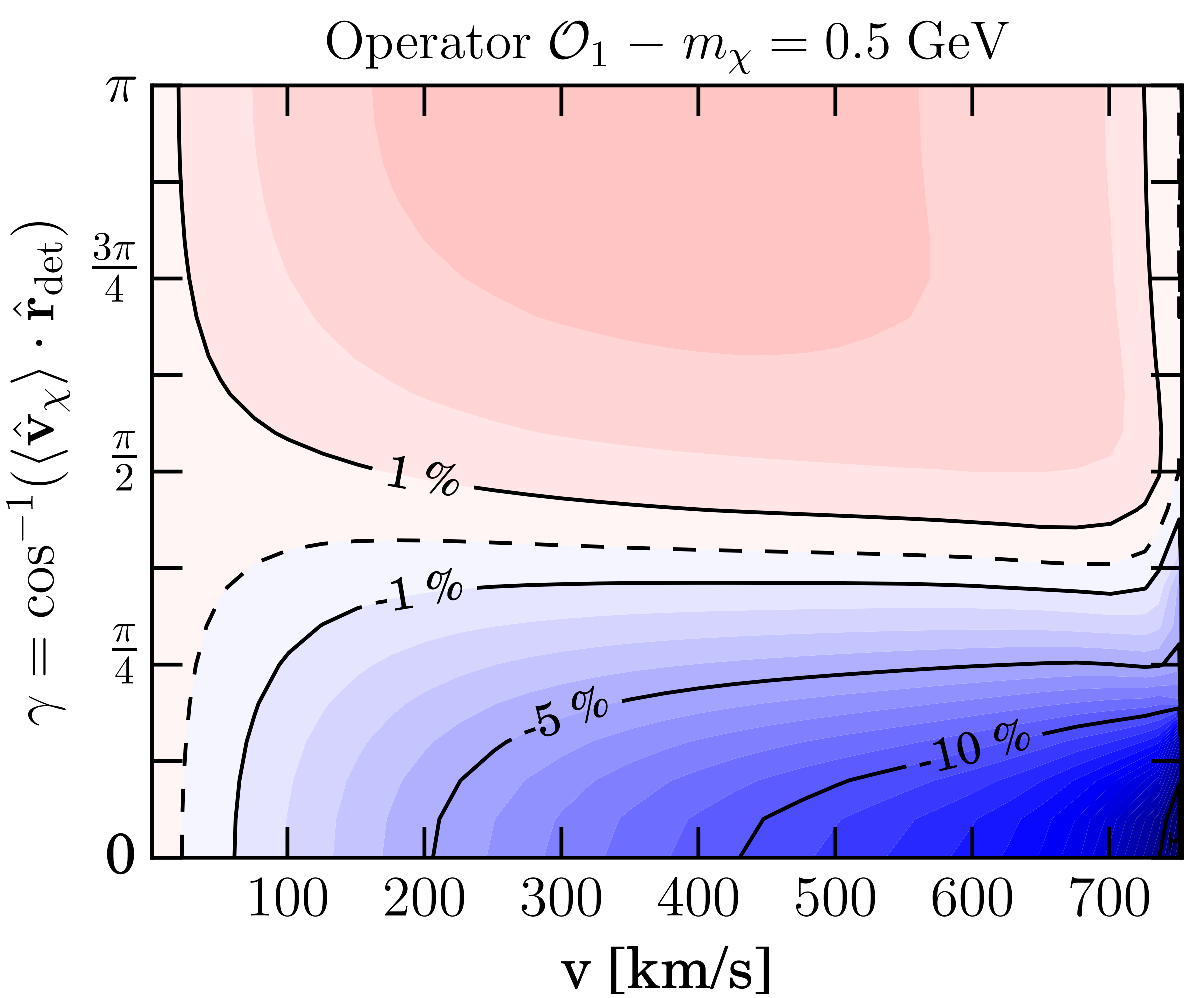
Calculate DM speed distribution after Earth scattering: $f_{\rm pert}(v)$

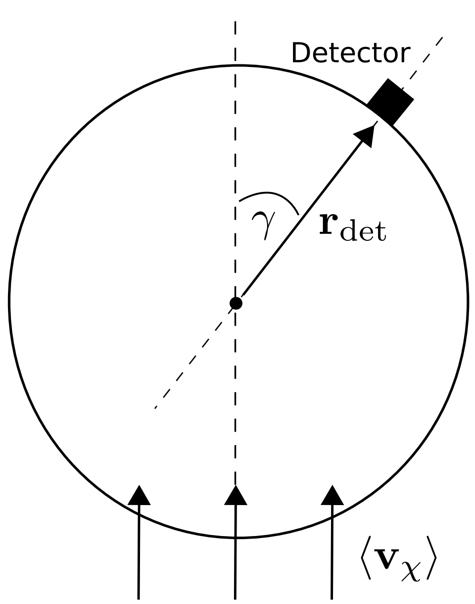


Percentage change in speed dist.

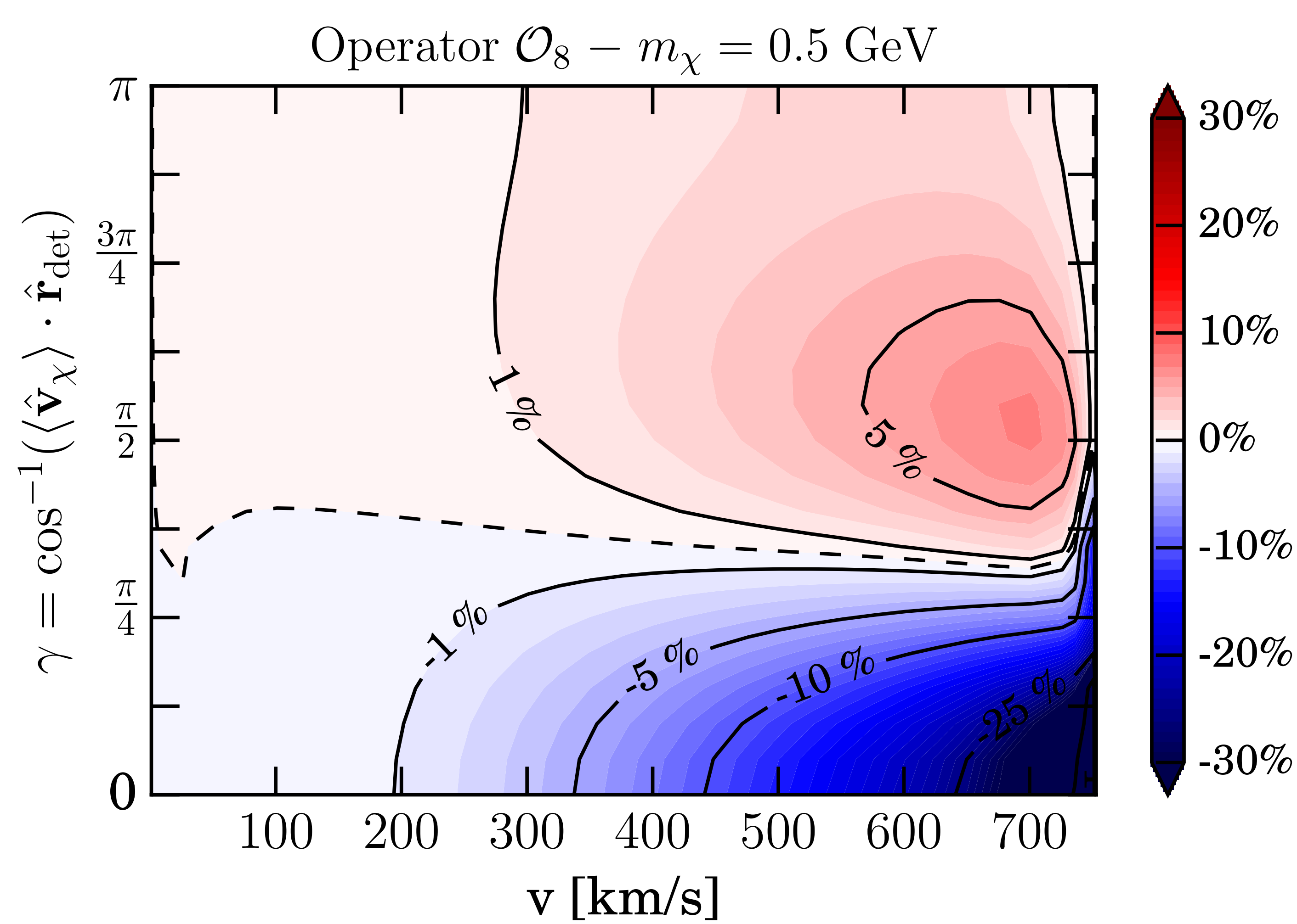


Speed Distribution - O_1 vs O_8

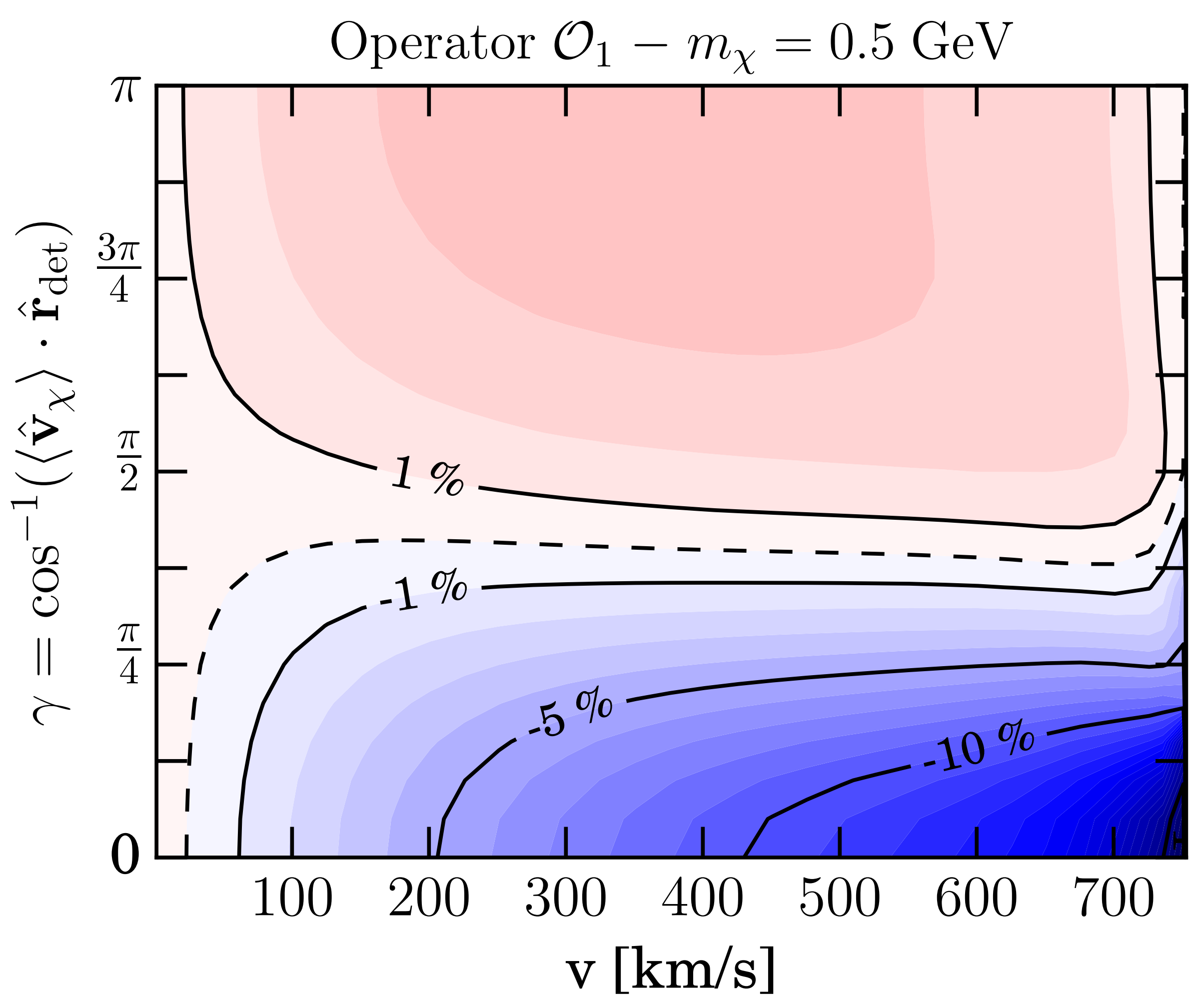




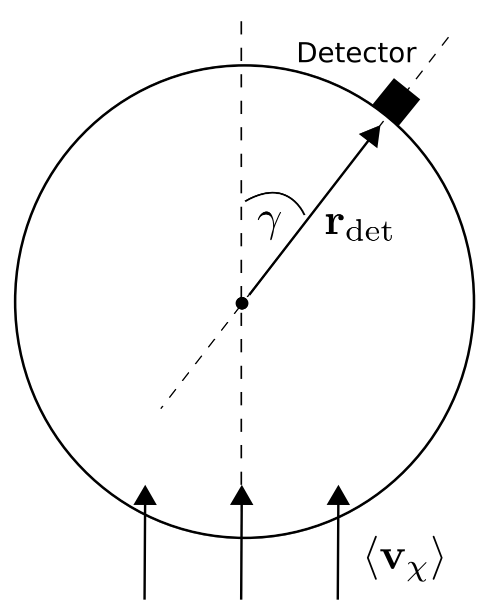
Operator 8 -
preferentially *forward* deflection

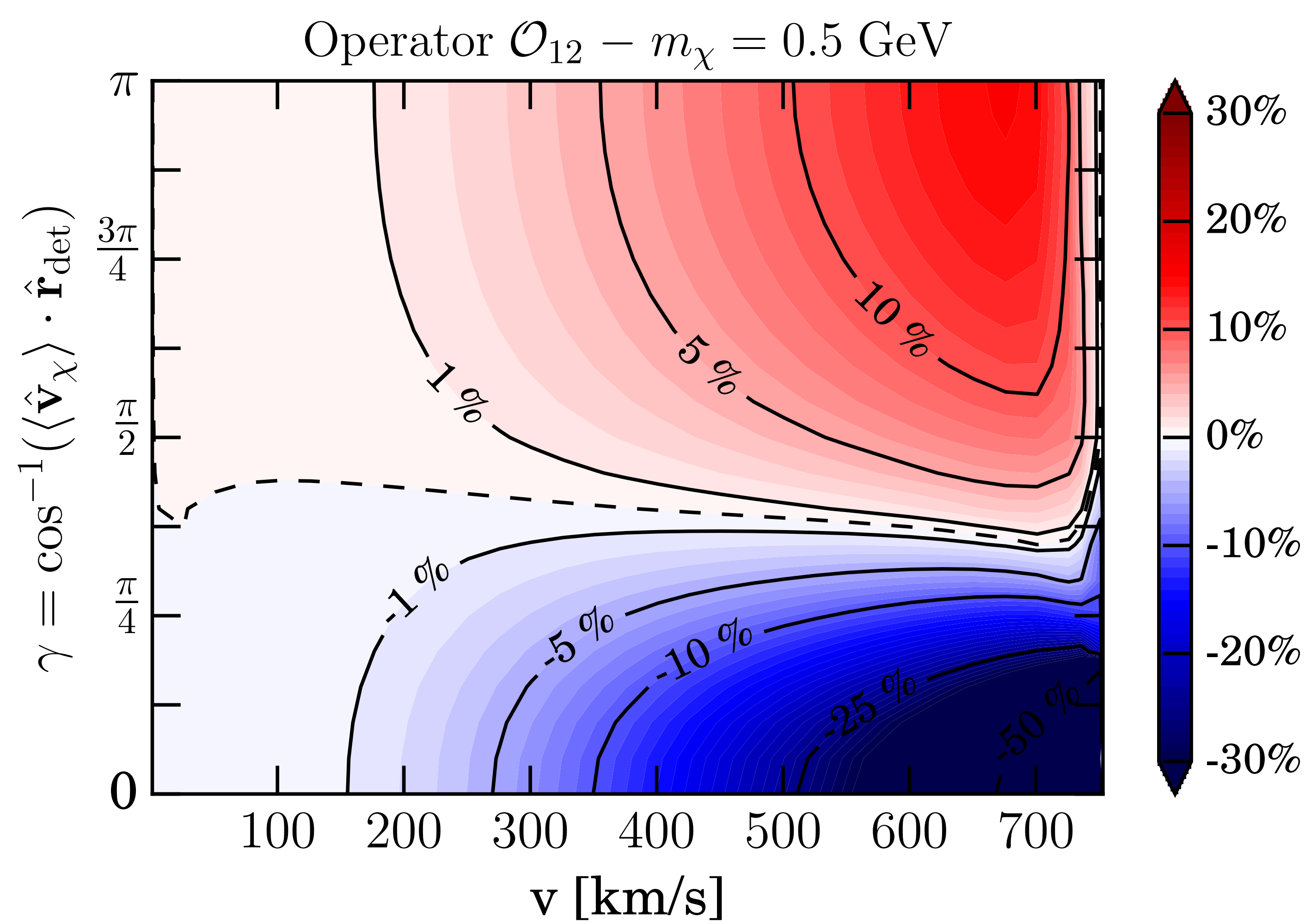


Speed Distribution - O_1 vs O_{12}

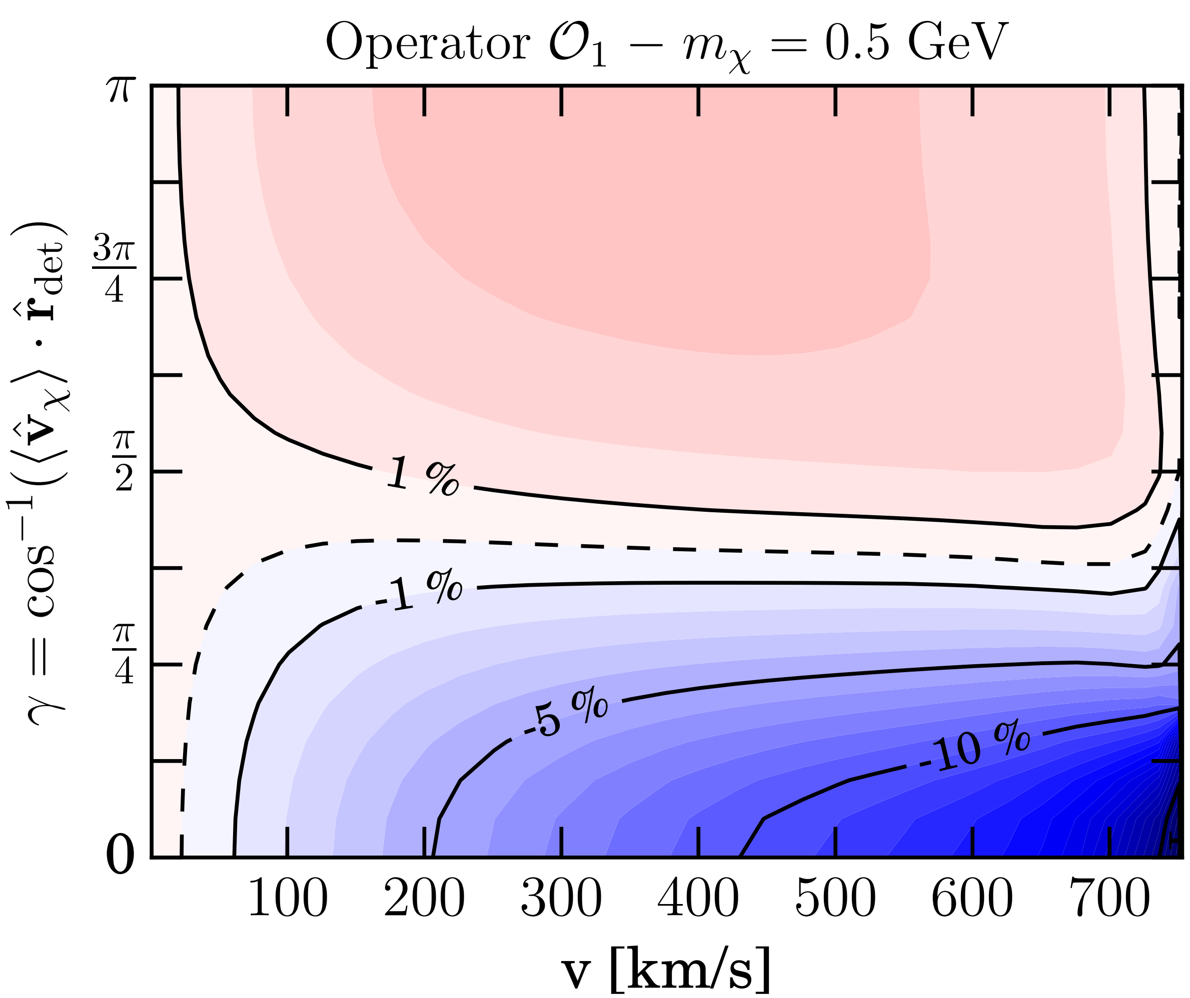


Operator 12 -
preferentially *backward* deflection

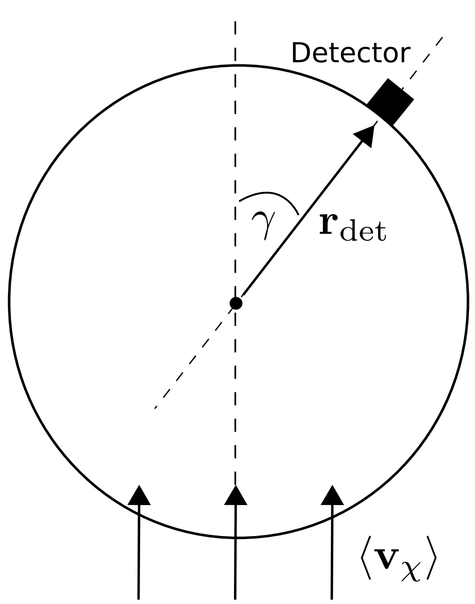


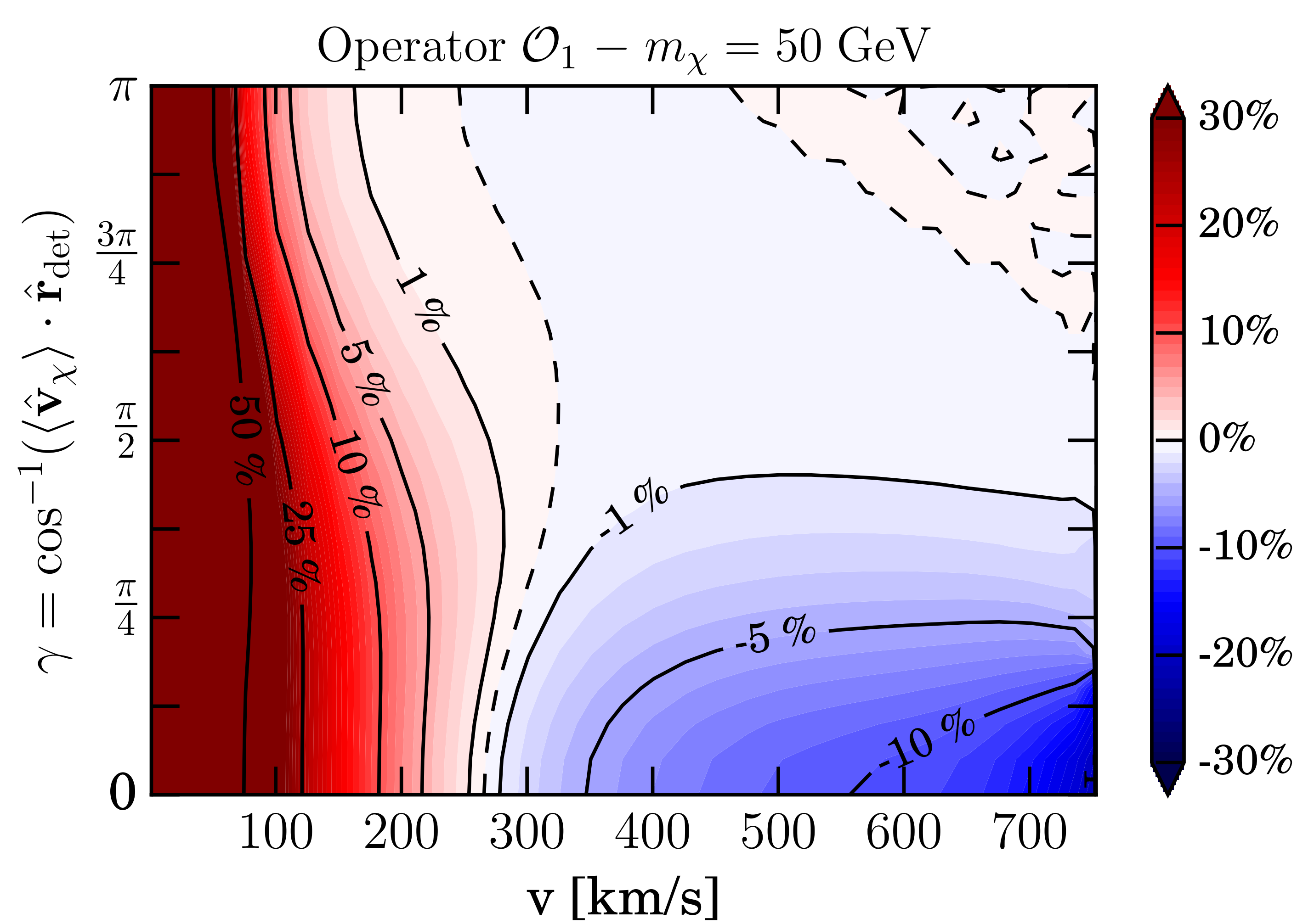


Low mass vs High mass



Higher mass DM





Sanity check

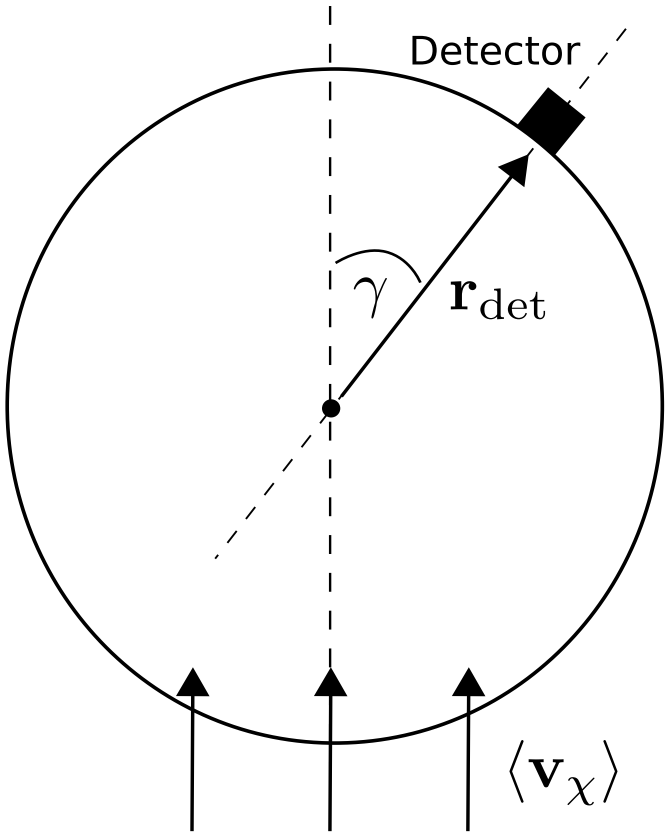
Compare rate of DM particles entering the Earth…

$$\Gamma_{\mathrm{in}} = \pi R_\oplus \langle v \rangle$$

…and rate of DM particle leaving the Earth…

$$\Gamma_{\mathrm{out}} = \int_{\mathbf{v}\cdot\mathbf{r}>0} \mathrm{d}^2\mathbf{r} \int \mathrm{d}^3\mathbf{v}\, f_{\mathrm{pert}}(\mathbf{v},\mathbf{r})\,(\mathbf{v}\cdot\mathbf{r})$$

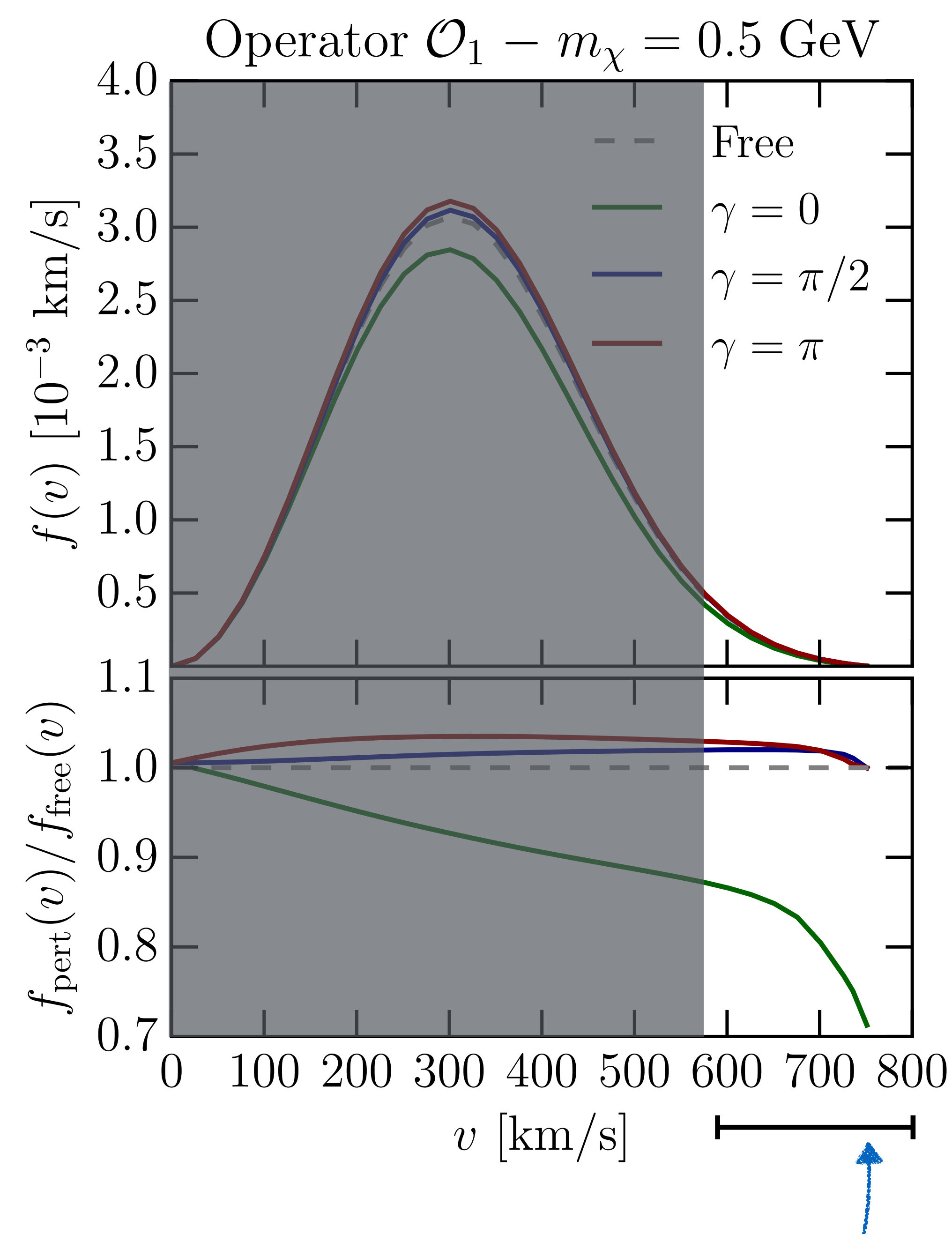
DM mass [GeV]	Operator	$\Delta\Gamma_{\mathrm{out}}^{\mathrm{Atten.}}/\Gamma_{\mathrm{in}}$	$\Delta\Gamma_{\mathrm{out}}^{\mathrm{Defl.}}/\Gamma_{\mathrm{in}}$	$\Gamma_{\mathrm{out}}/\Gamma_{\mathrm{in}}$
0.5	$\hat{\mathcal{O}}_1$	-7.8%	$+7.0\%$	99.2%
0.5	$\hat{\mathcal{O}}_8$	-8.0%	$+7.3\%$	99.2%
0.5	$\hat{\mathcal{O}}_{12}$	-7.8%	$+7.2\%$	99.4%
50	$\hat{\mathcal{O}}_1$	-7.5%	$+7.3\%$	99.9%
50	$\hat{\mathcal{O}}_8$	-8.0%	$+8.4\%$	100.4%
50	$\hat{\mathcal{O}}_{12}$	-7.3%	$+6.6\%$	99.3%



Event Rate

Calculate number of signal events in a CRESST-II like experiment, with and without the effects of Earth-Shadowing, N_{pert} and N_{free} .

Scattering predominantly with Oxygen and Calcium.



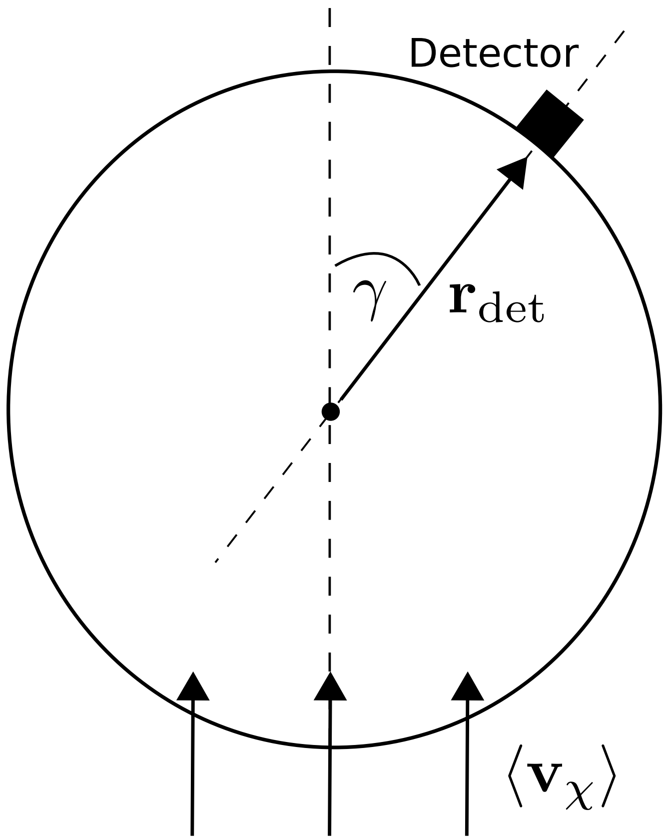
DM particles within $3\,\sigma_E$ of the energy threshold
$E_{\mathrm{th}} \sim 300$ eV

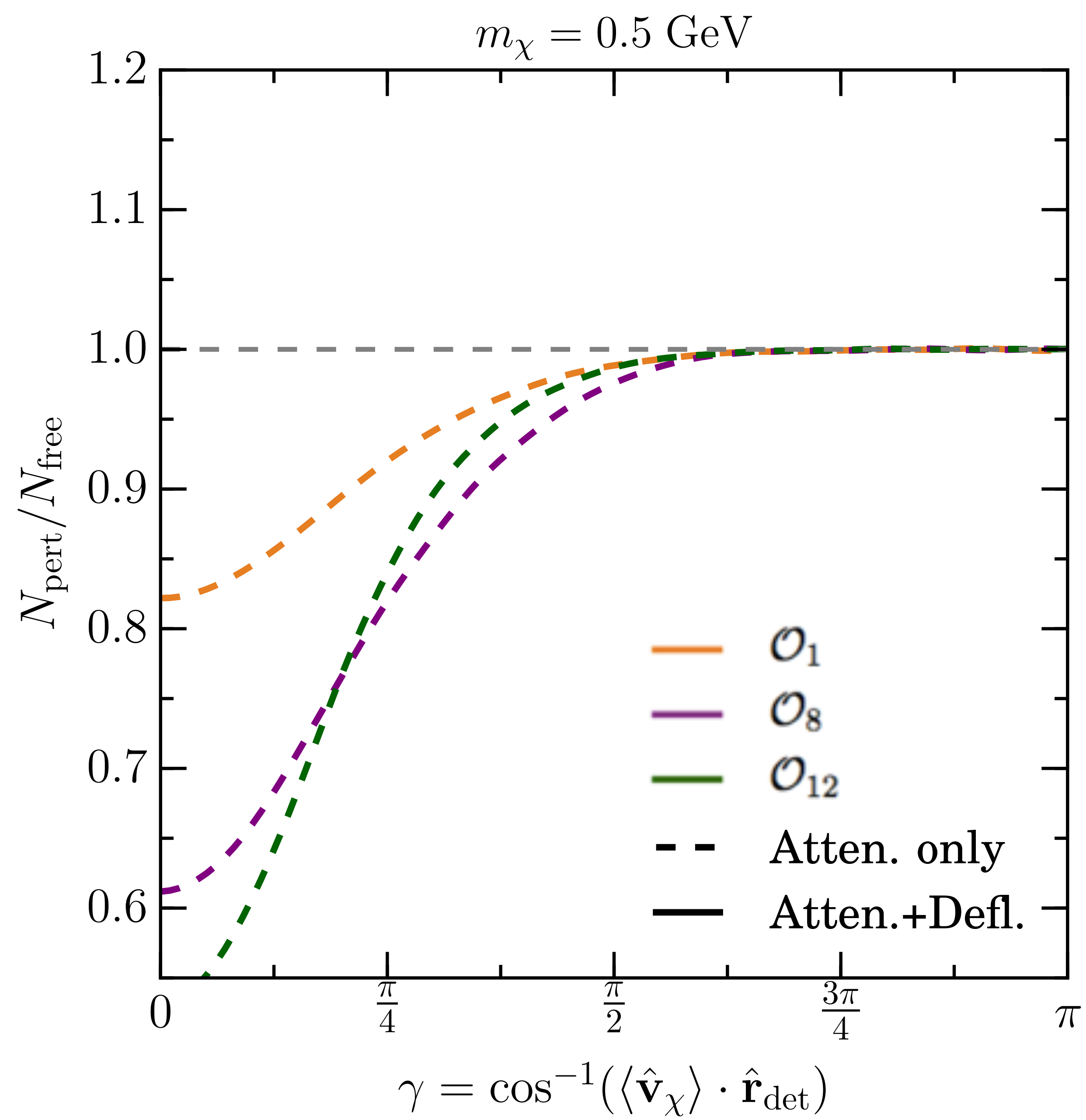
CRESST-II Rate (attenuation-only)

Operator 1 - isotropic deflection

Operator 8 - forward deflection

Operator 12 - backward deflection



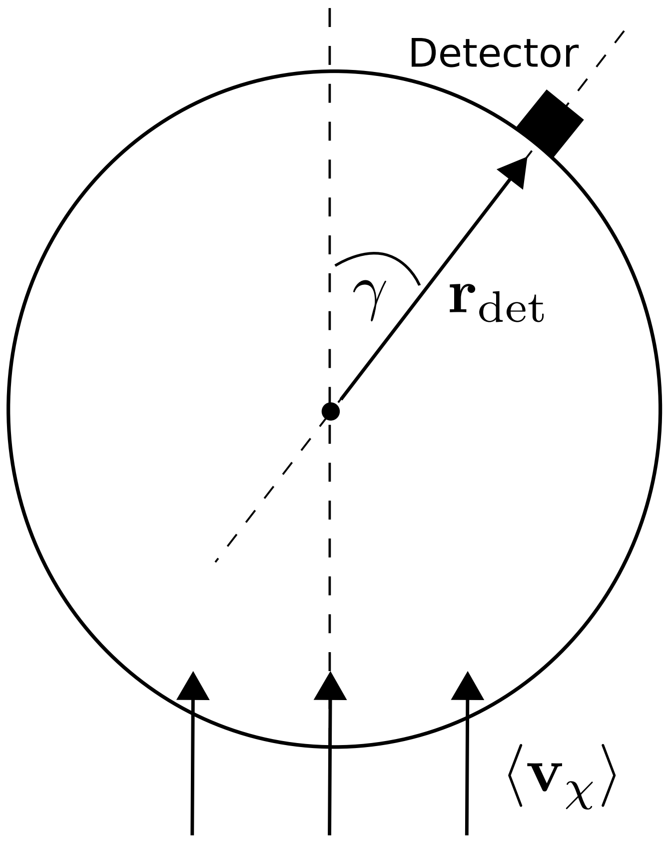


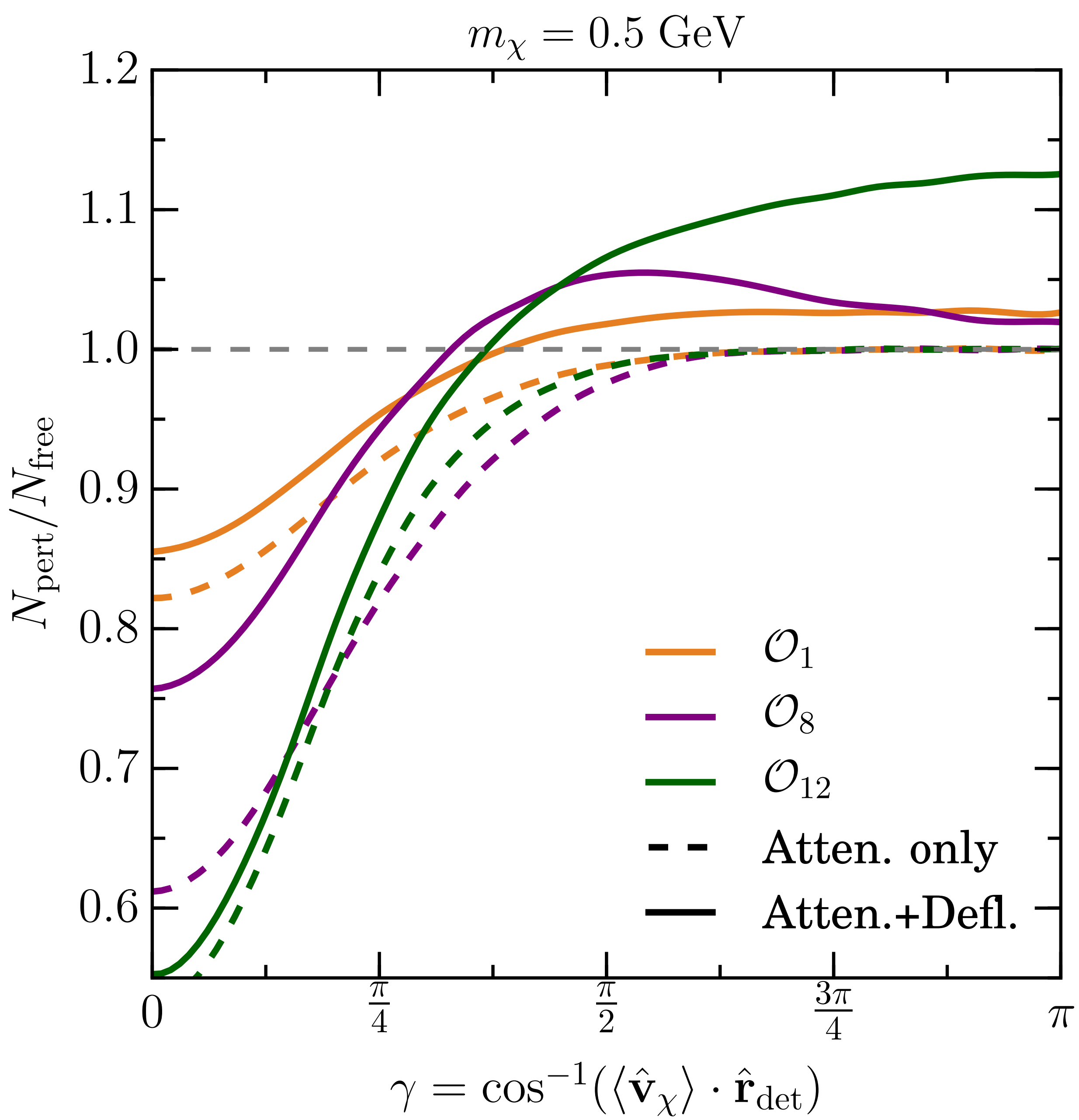
CRESST-II Rate (attenuation + deflection)

Operator 1 - isotropic deflection

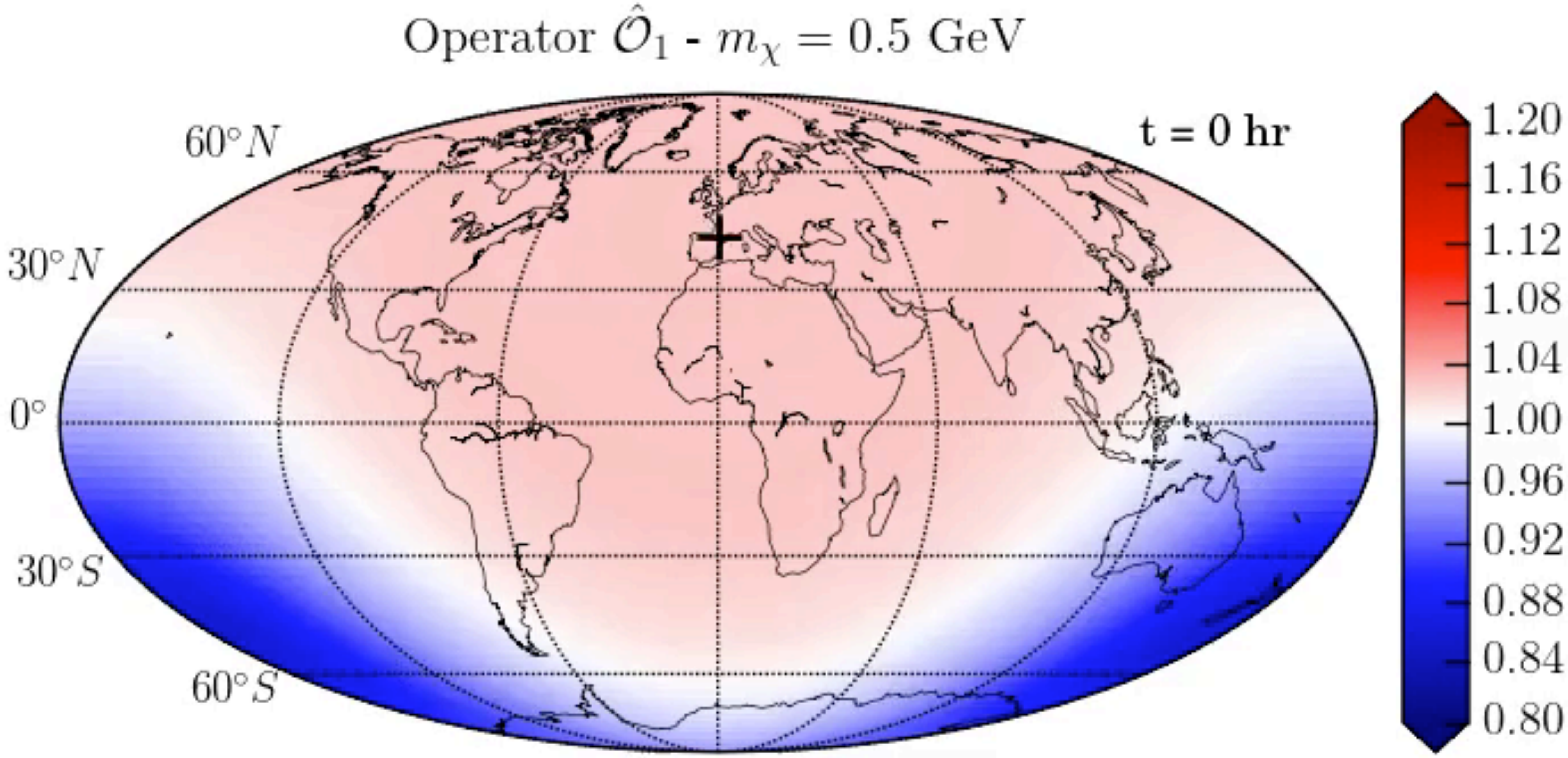
Operator 8 - forward deflection

Operator 12 - backward deflection





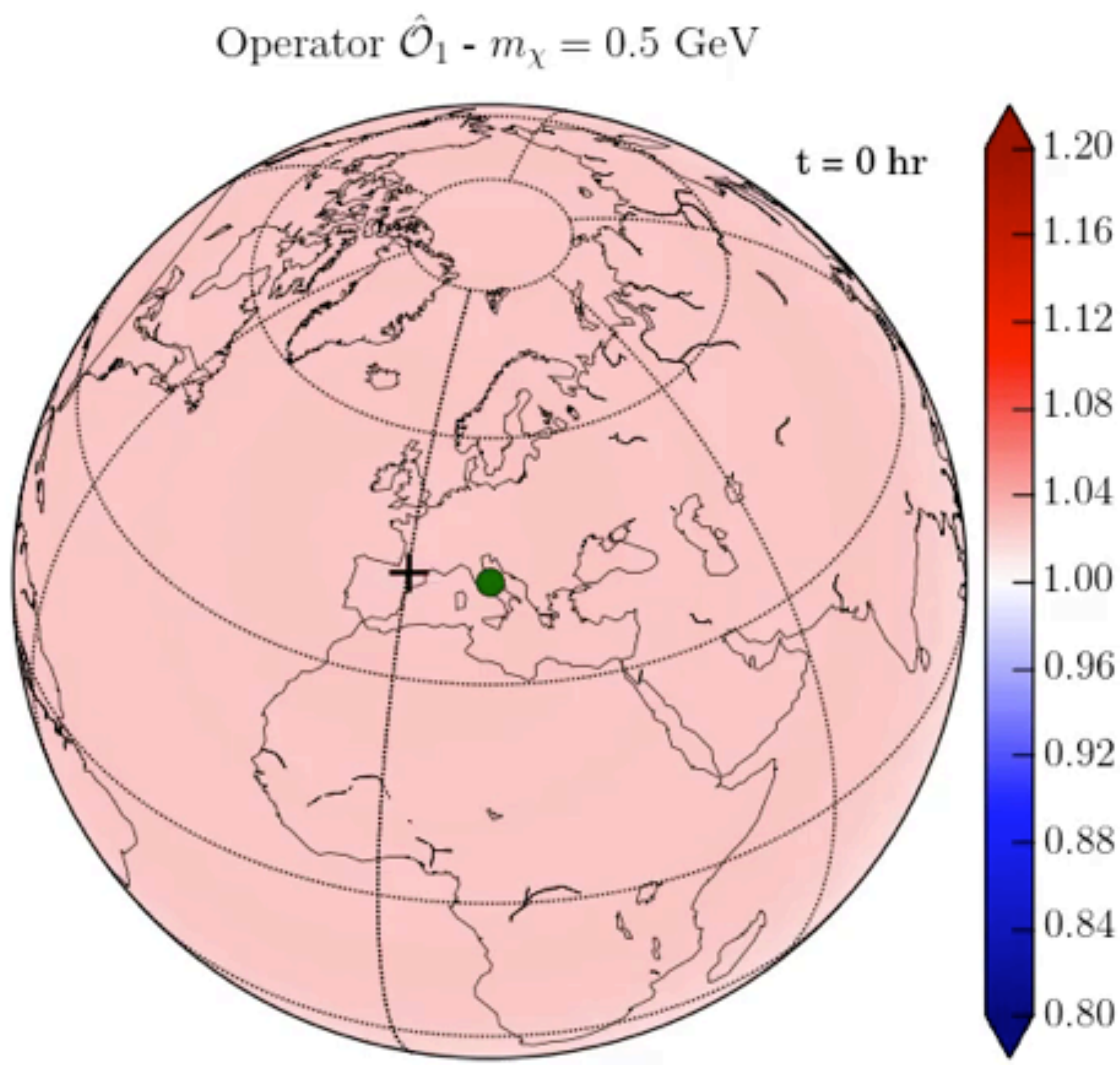
Mapping the CRESST-II Rate



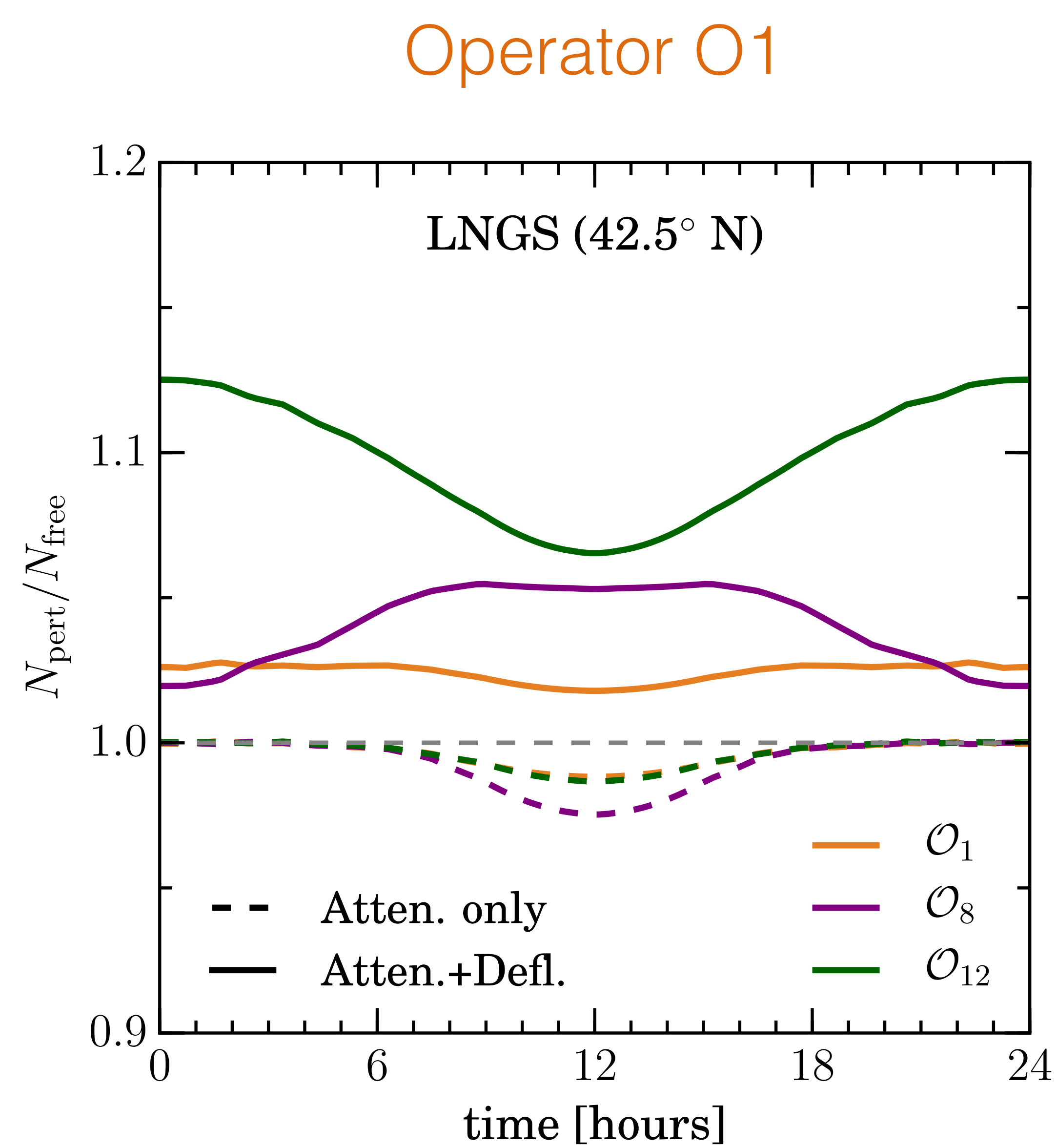
Relative rate enhancement due to Earth-scattering *(attenuation + deflection)*

LNGS - Operator 1

LNGS - Gran Sasso Lab, Italy

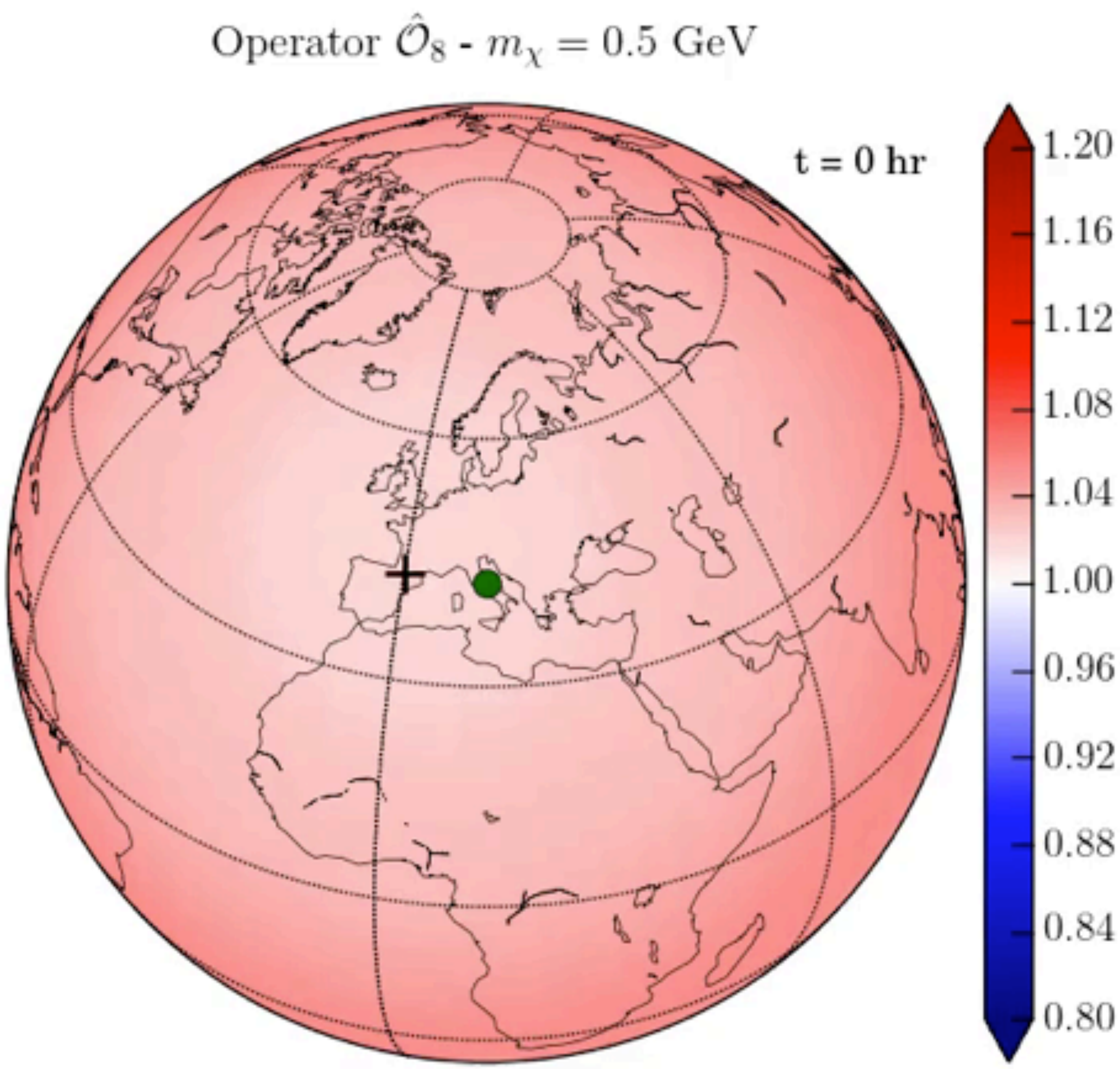


Relative rate enhancement due to Earth-scattering *(attenuation + deflection)*



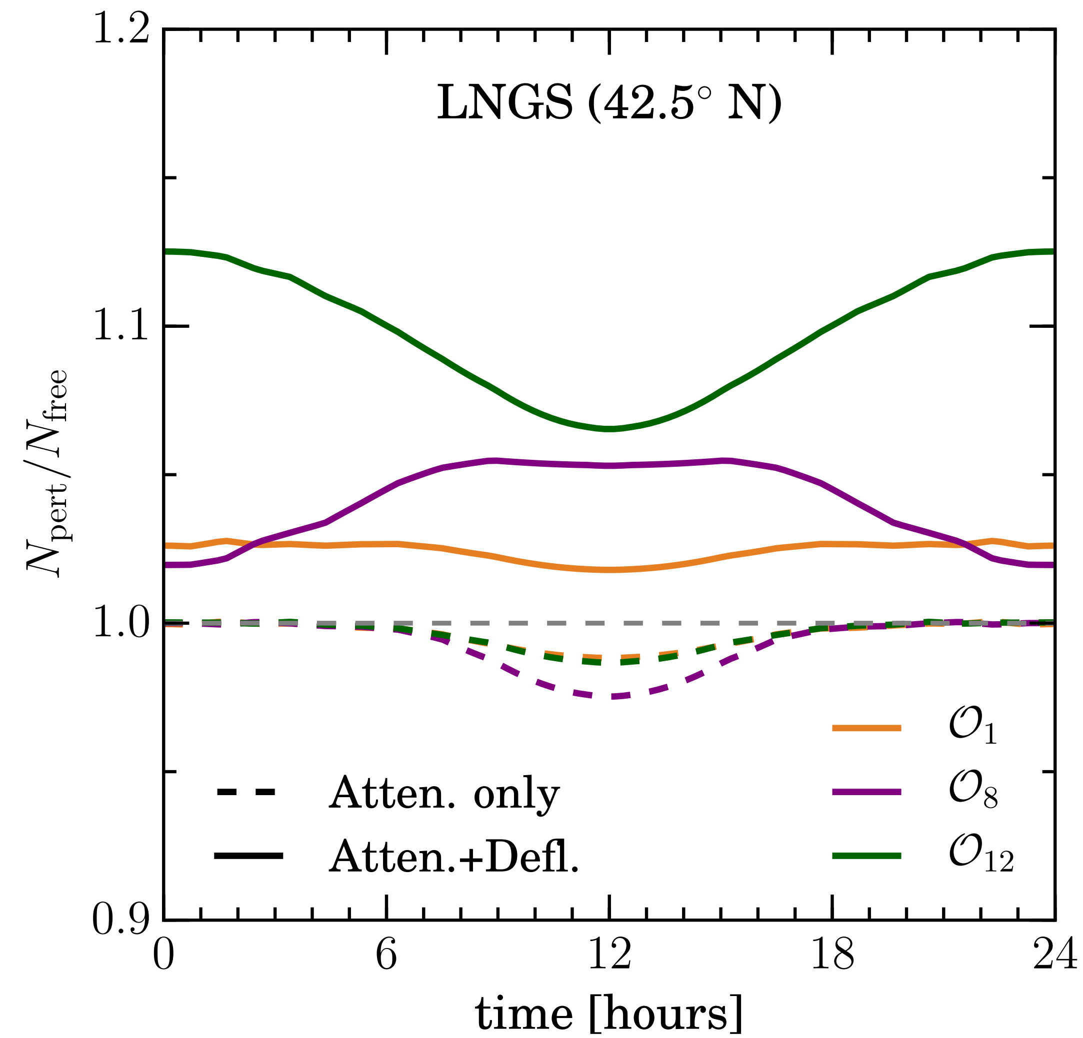
LNGS - Operator 8

LNGS - Gran Sasso Lab, Italy



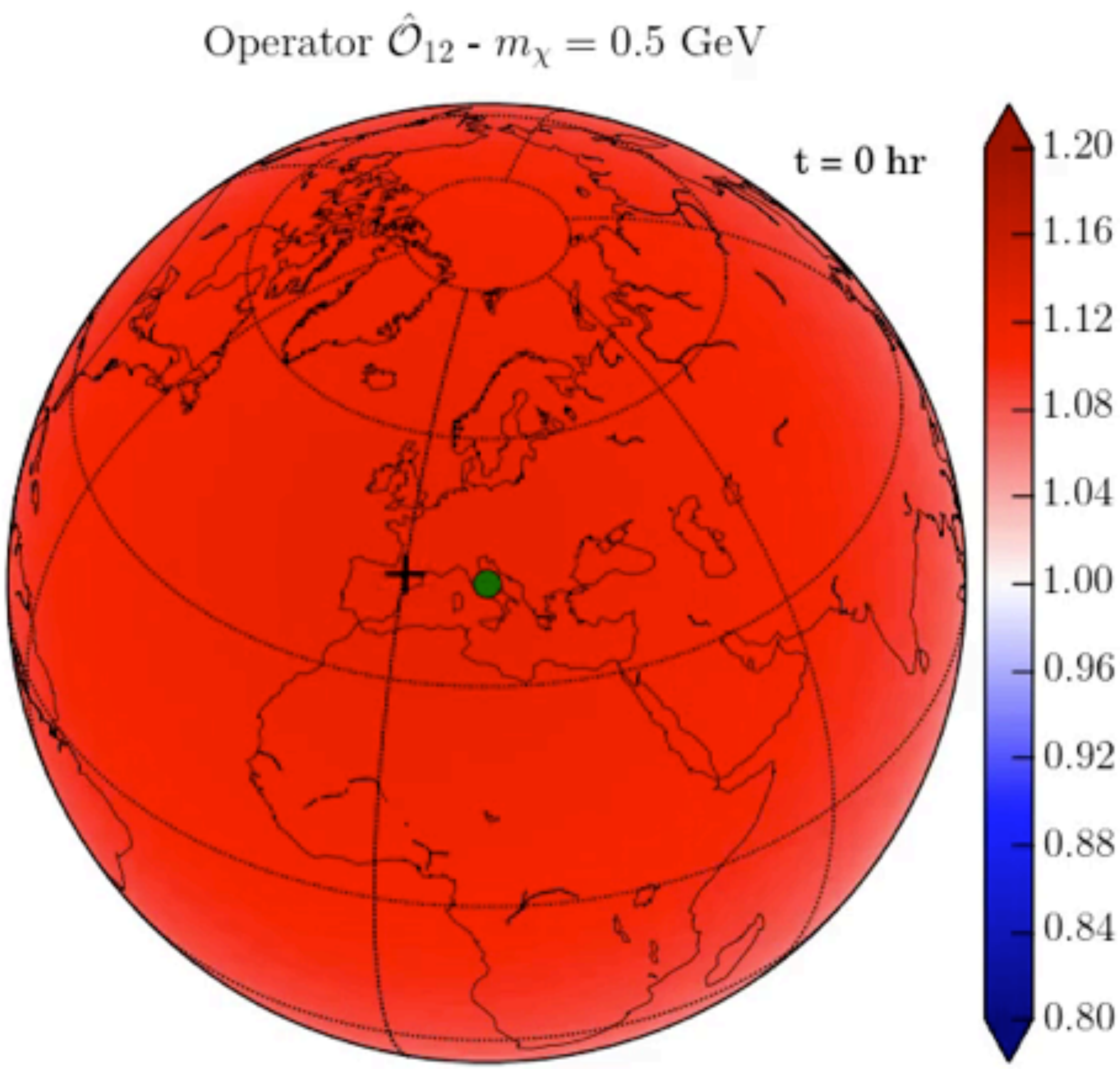
Relative rate enhancement due to Earth-scattering *(attenuation + deflection)*

Operator O8



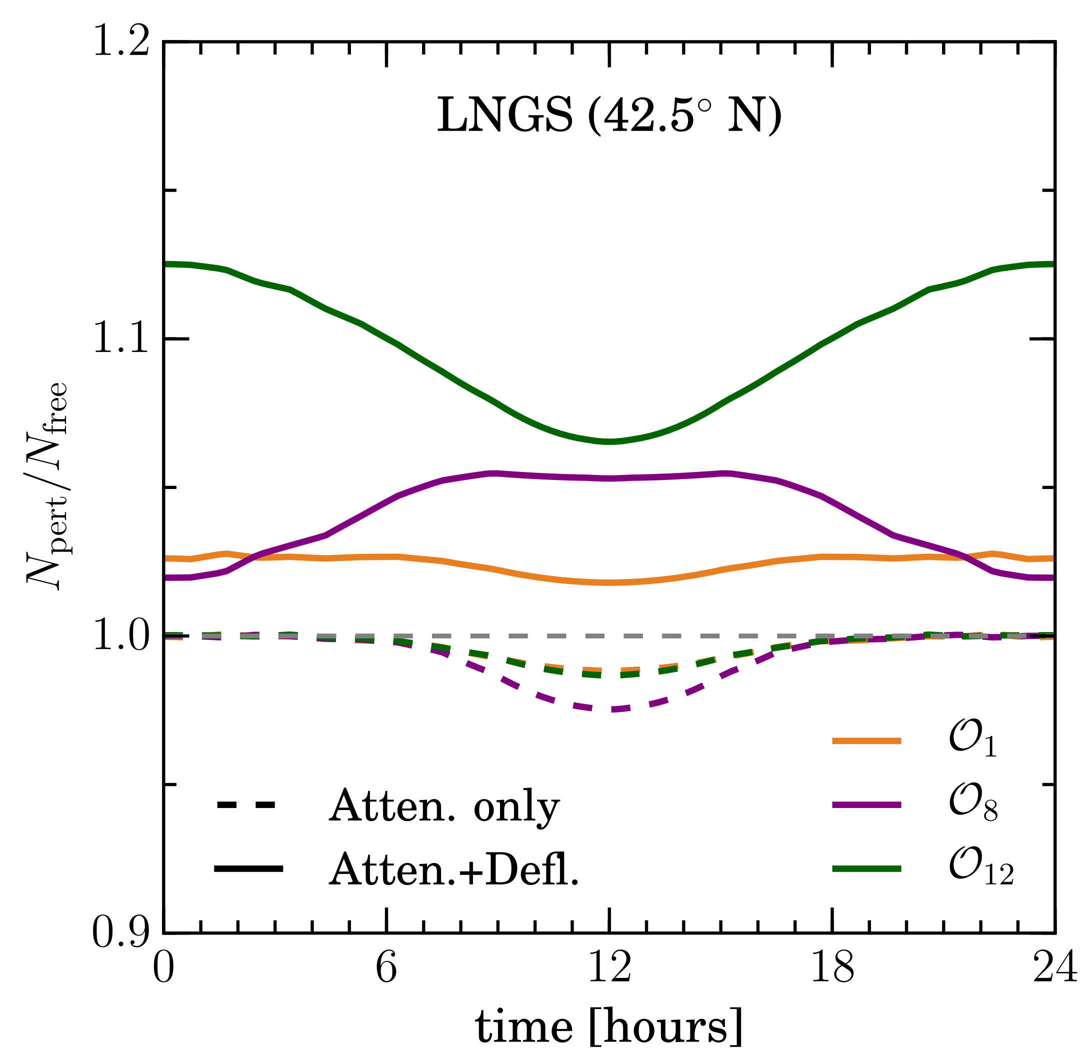
LNGS - Operator 12

LNGS - Gran Sasso Lab, Italy



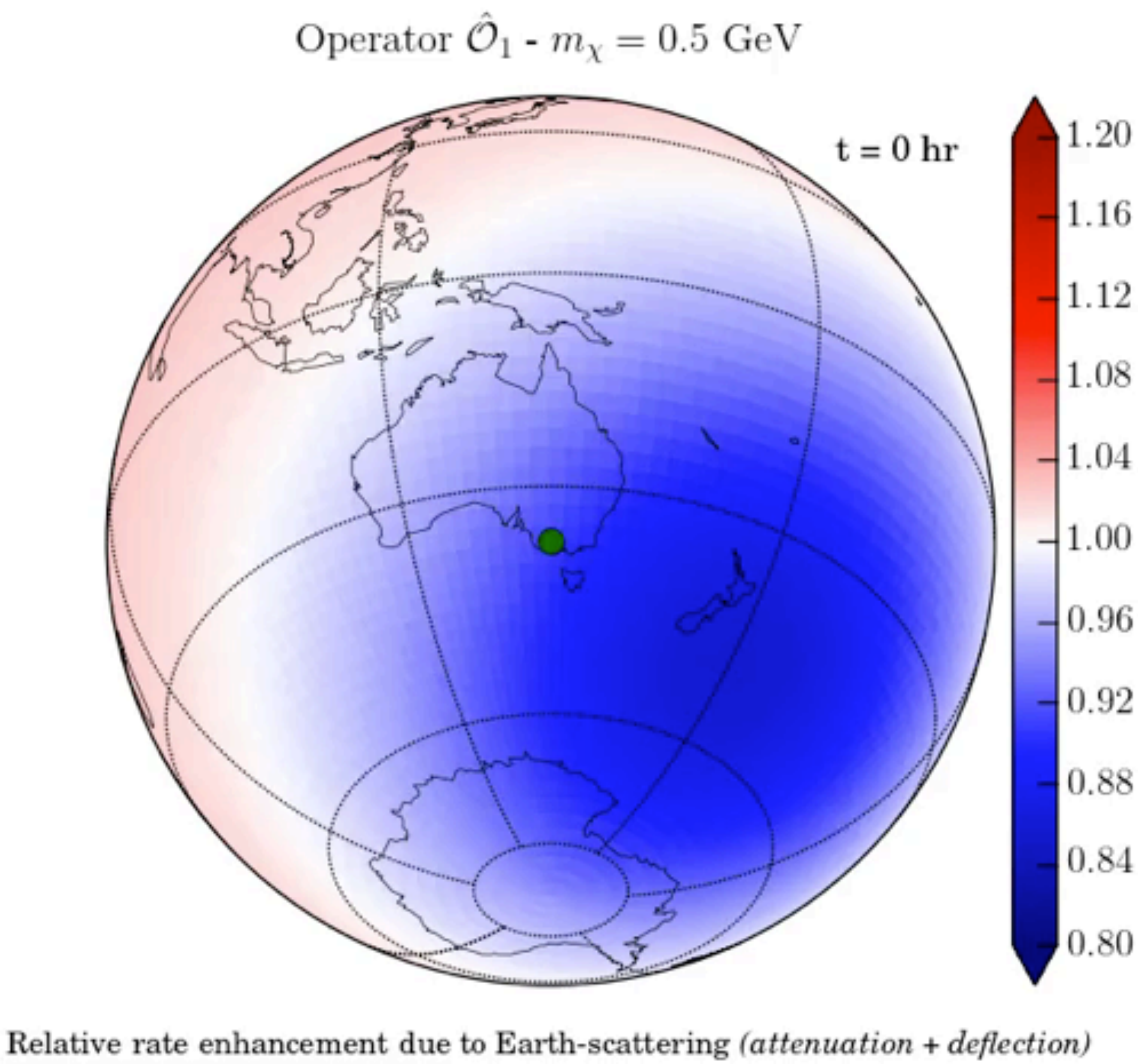
Relative rate enhancement due to Earth-scattering *(attenuation + deflection)*

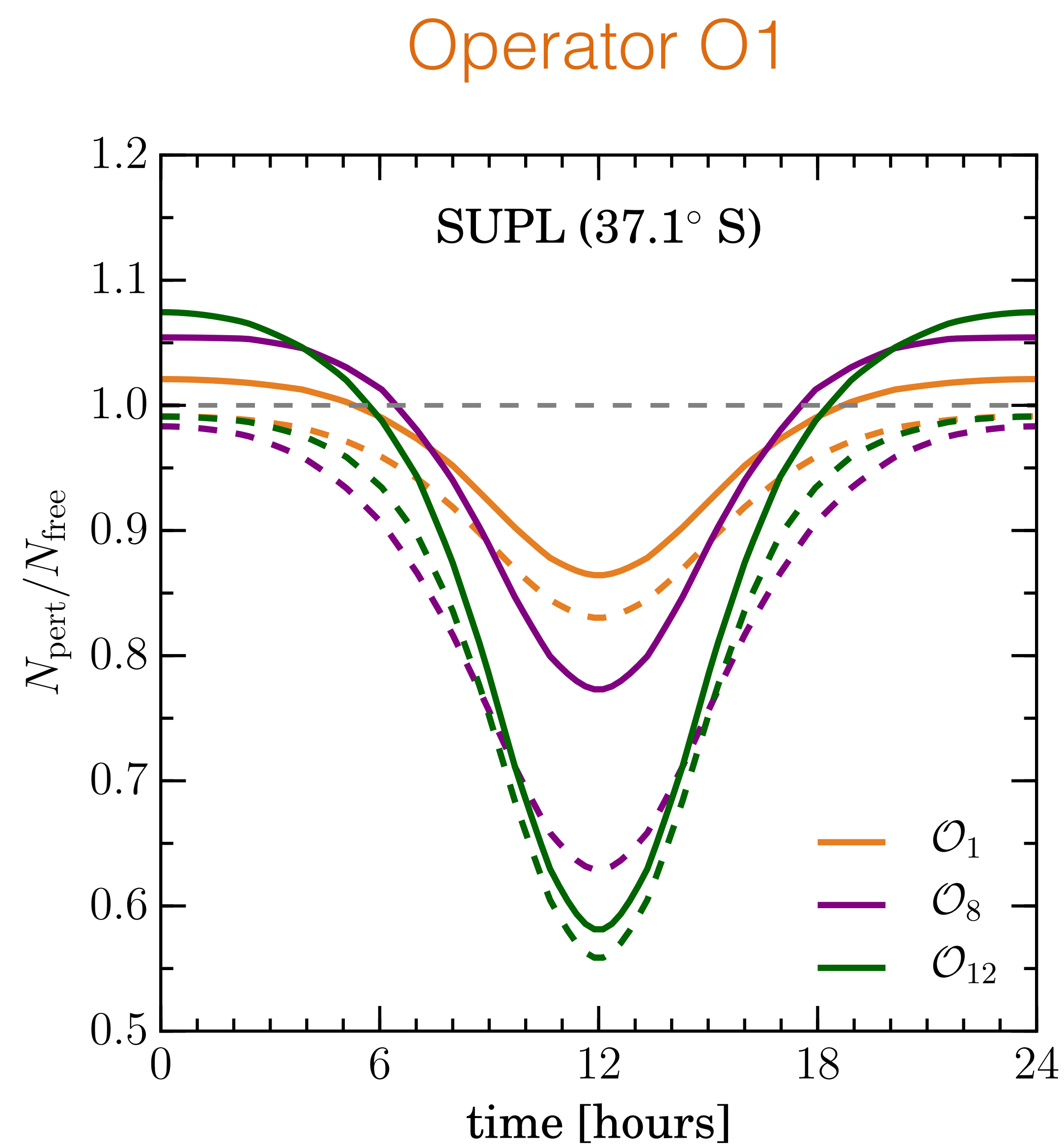
Operator O12



SUPL - Operator 1

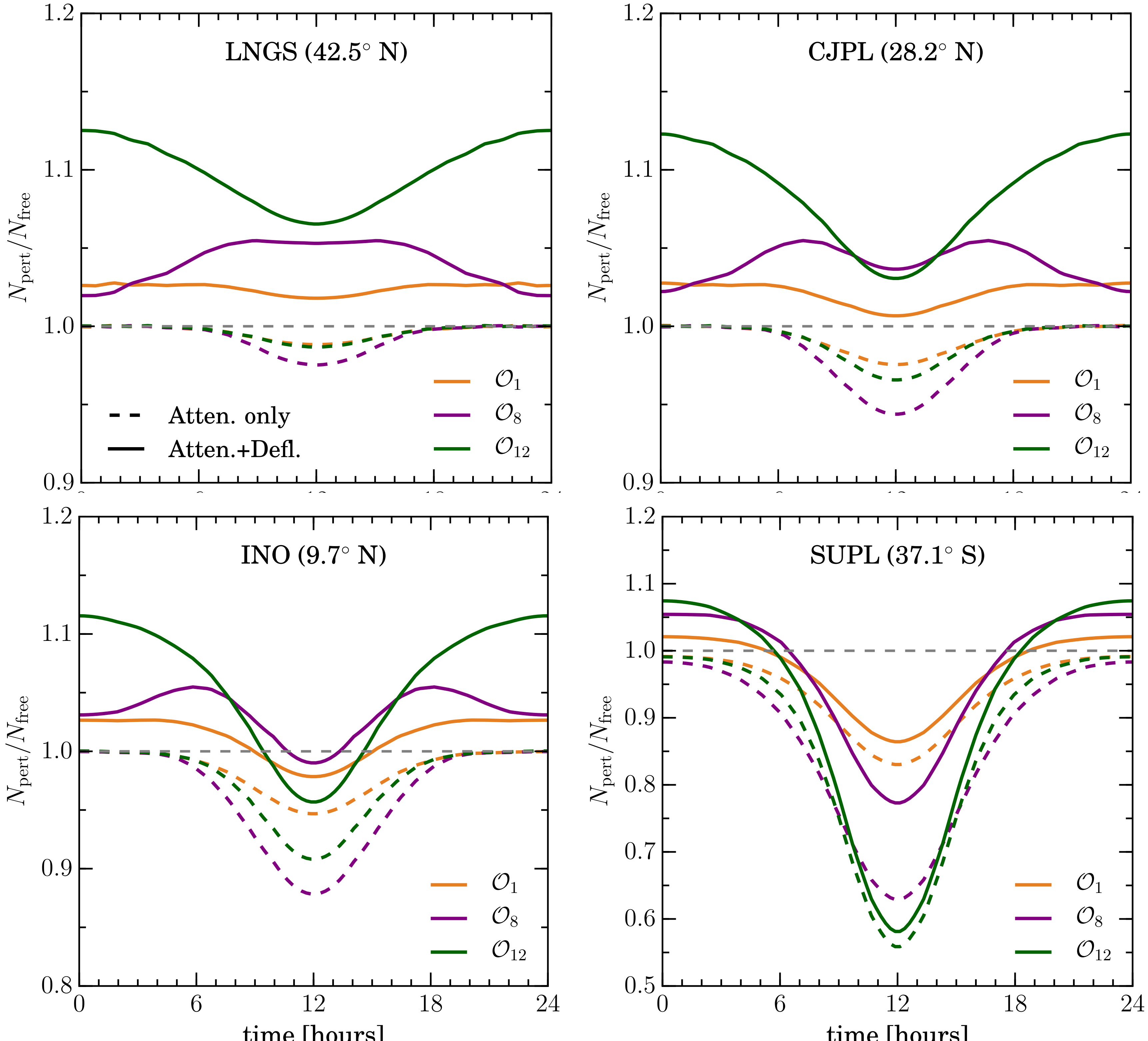
SUPL - Stawell Underground Physics Lab, Australia





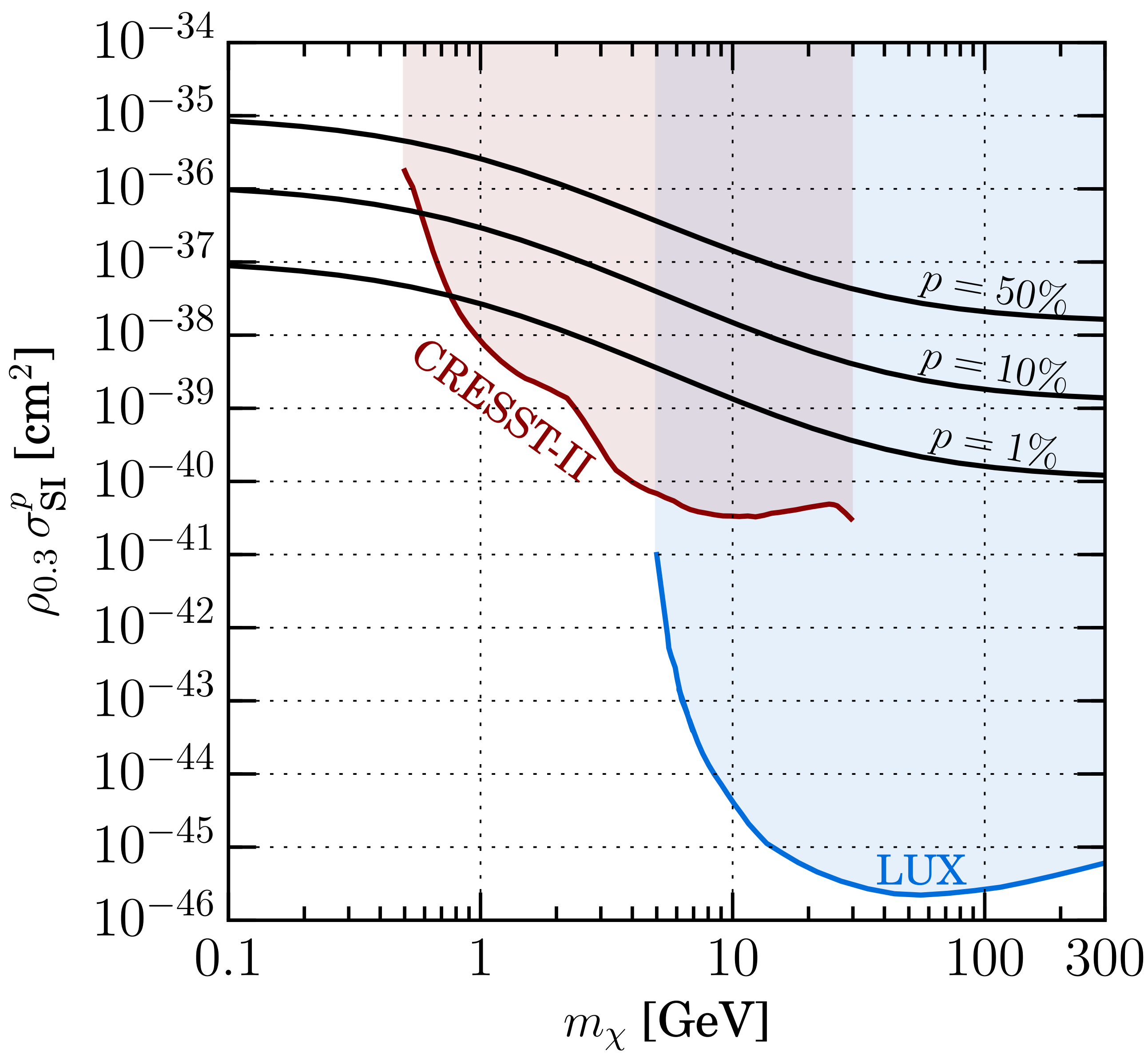
Around the world

China Jinping Lab



India-based
Neutrino Observatory

Implications of Earth-Shadowing



Smoking gun signature:
daily modulation +
location dependence

Possibility to distinguish different interactions with distinctive modulation signals

Possibility to measure the local DM density (by breaking degeneracy with cross section)

Future work

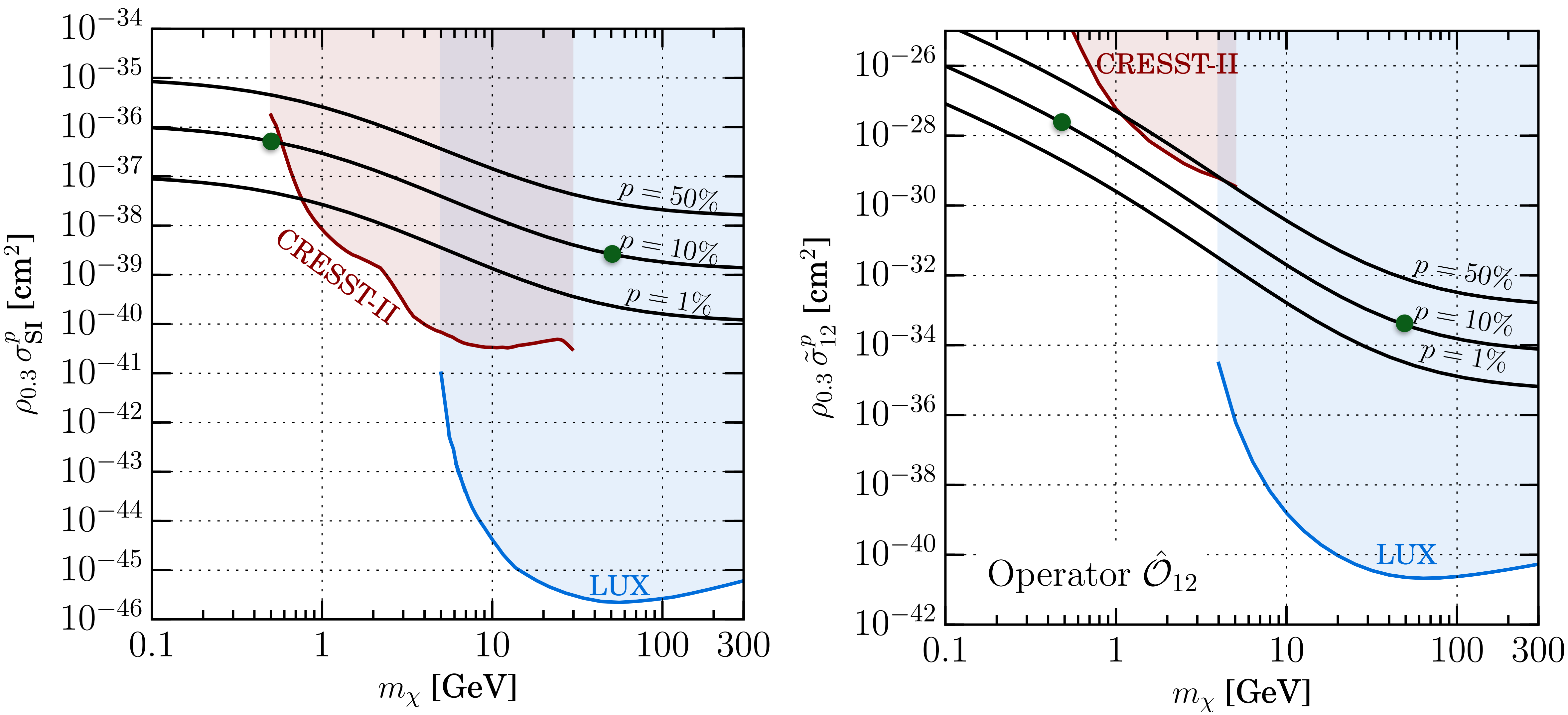
Here, we have considered only the DM *speed* distribution. Need to look at the full 3-D *velocity* distribution to explore directional signatures of Earth-Shadowing.

The Single-scatter approximation is important to capture the effects of deflection. But it will break down rapidly as we increase the DM cross section. Next steps:

- Calculations in the many-scatter/'diffusion' regime
- Dedicated simulations to test the single-scatter regime and connect to very high cross sections (work in progress by Chris Kouvaris)

Mapping out the parameter space

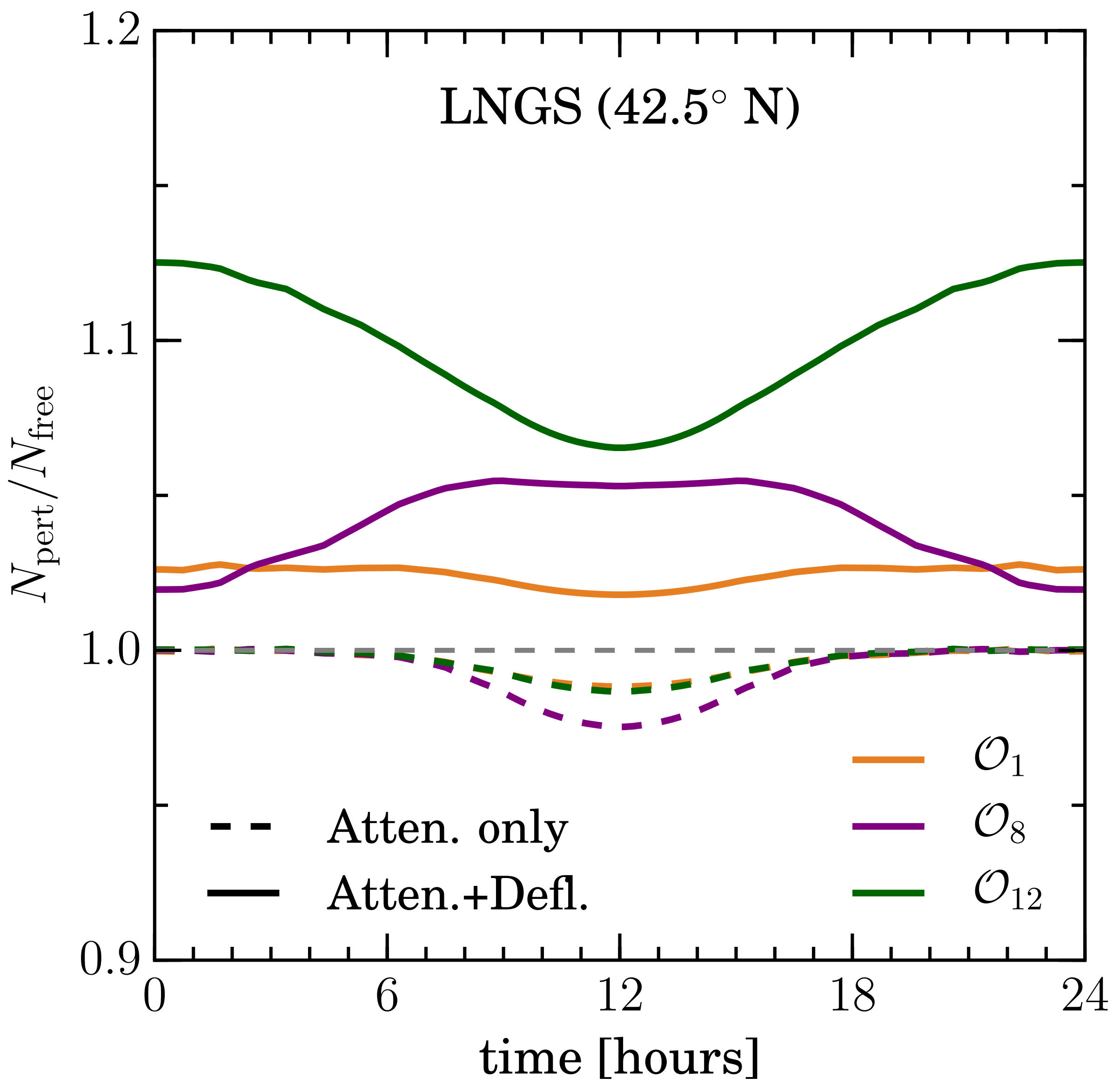
Continue mapping out parameter space (m_χ, σ_p) and explore impact on upper limits for a range of interactions…



…and encourage experimental collaborations to explore full NREFT parameter space.

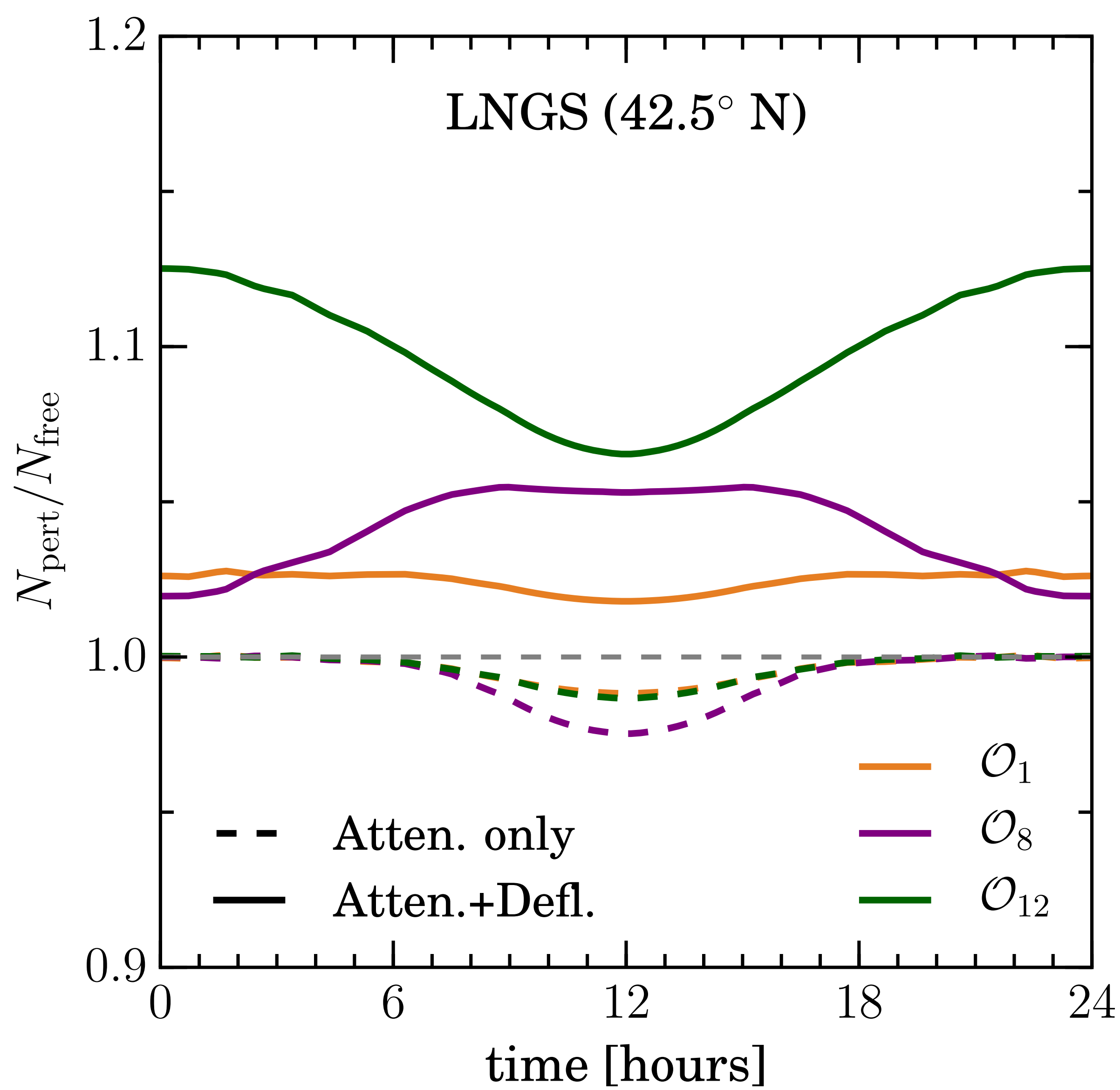
Conclusions

- Significant Earth-Shadowing is still allowed and detectable by current experiments
- Need to include both attenuation and deflection of DM
- Careful calculation including multiple elements, correct density profiles and different interactions



Conclusions

- Significant Earth-Shadowing is still allowed and detectable by current experiments
- Need to include both attenuation and deflection of DM
- Careful calculation including multiple elements, correct density profiles and different interactions
- The average incoming DM direction varies with time - distinctive daily modulation signals
- Different interactions may lead to modulations with different size and phases - and may therefore be distinguishable
- EARTHSHADOW code available online to include these effects: github.com/bradkav/EarthShadow



Conclusions

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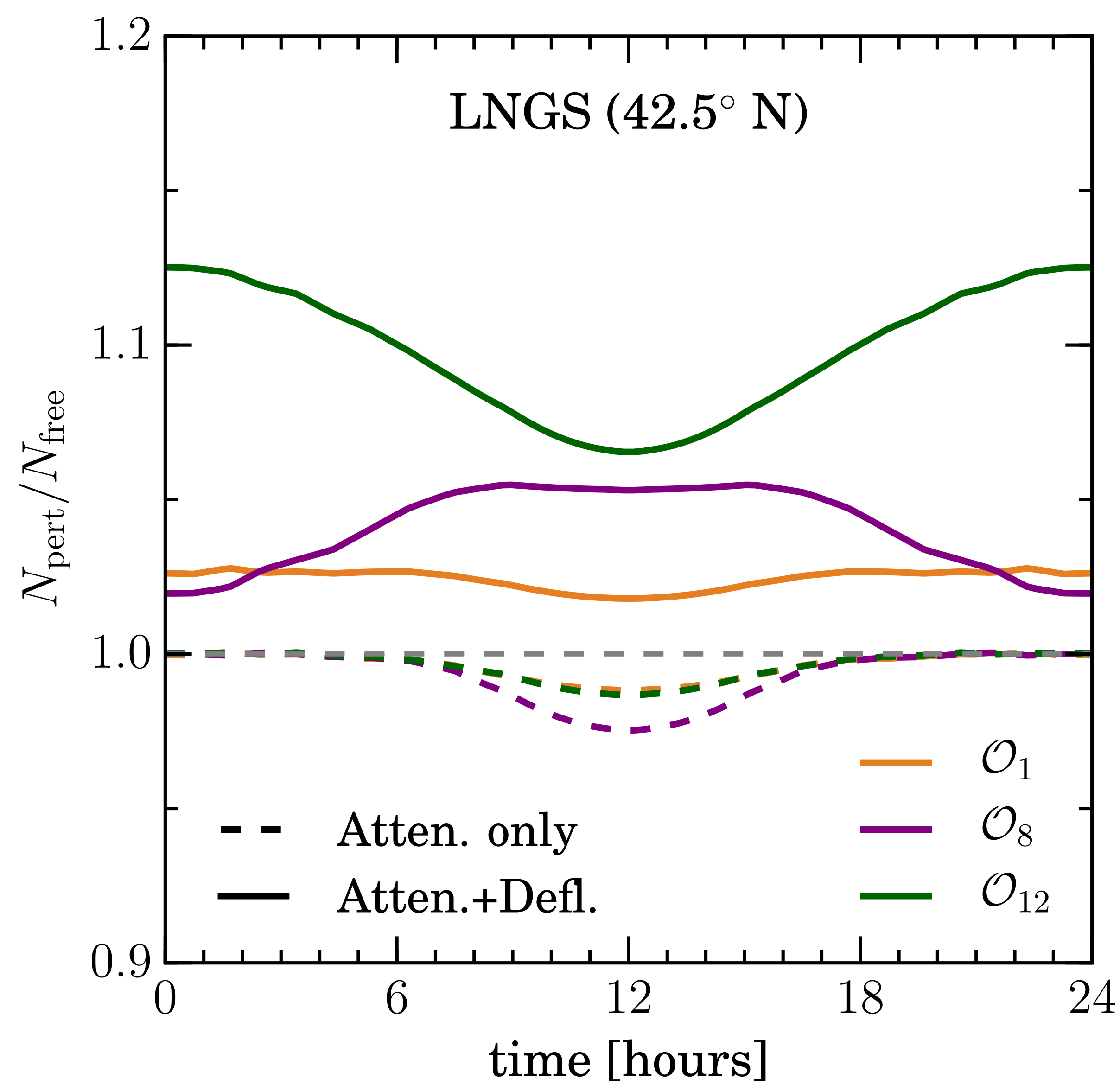
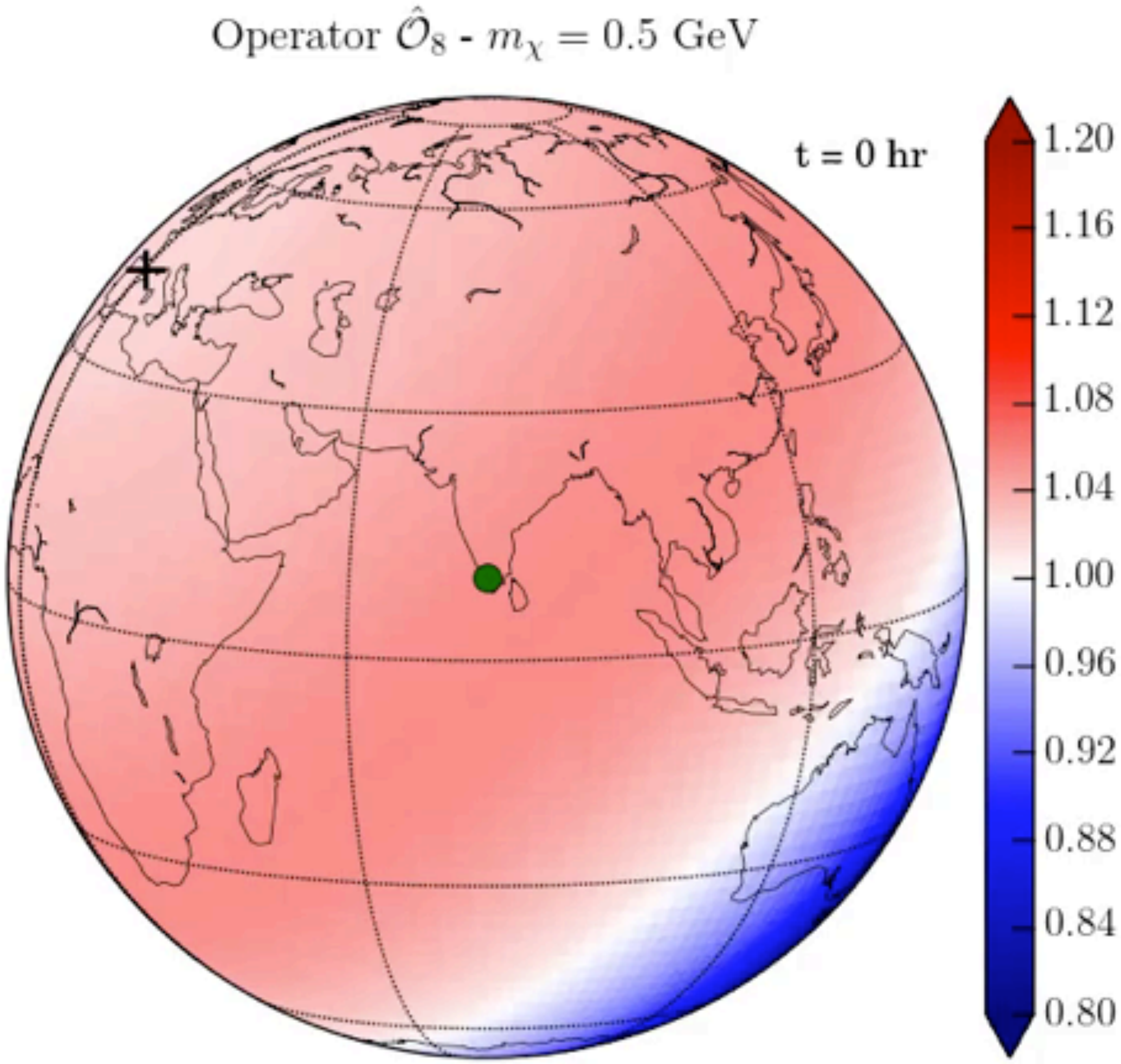


Thank you!

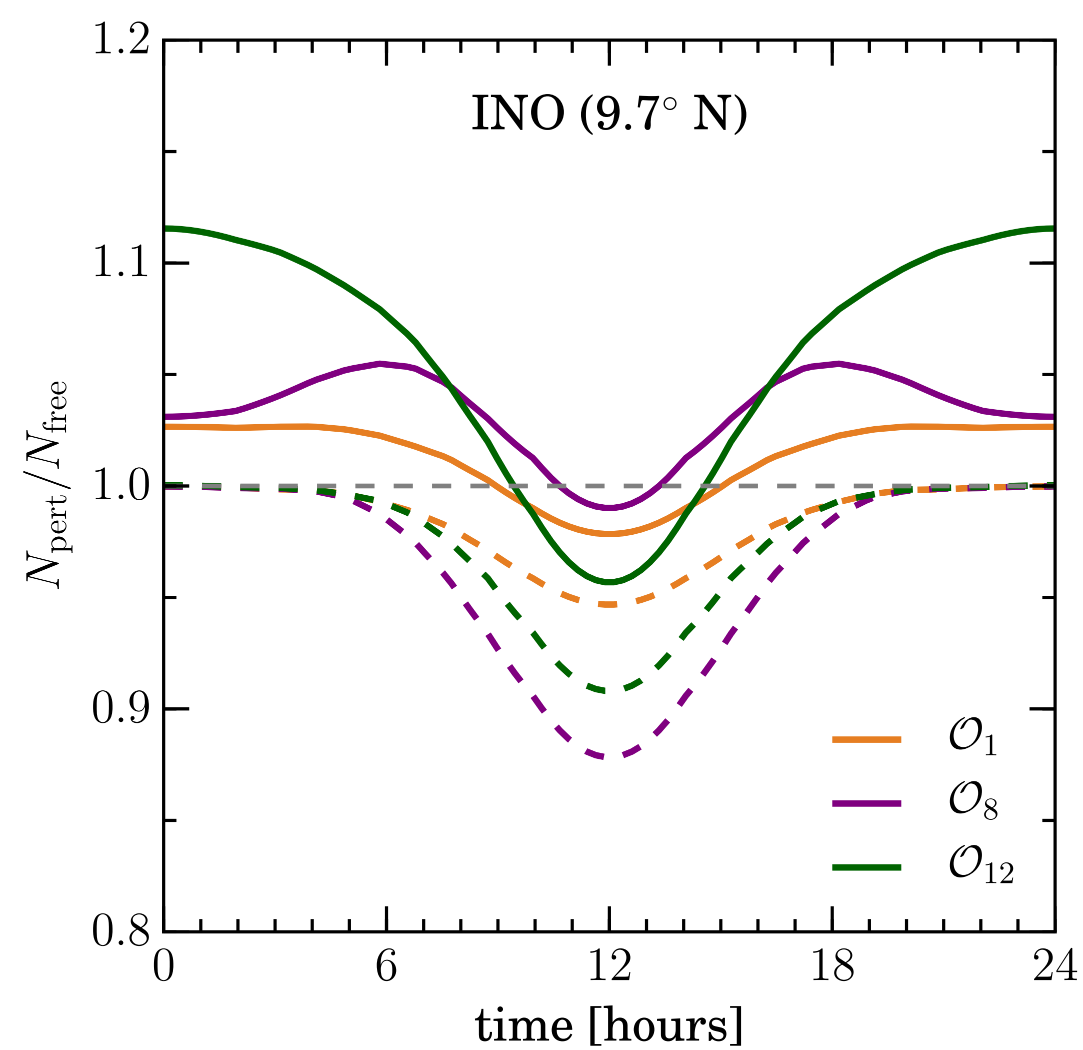
Backup Slides

INO - Operator 8

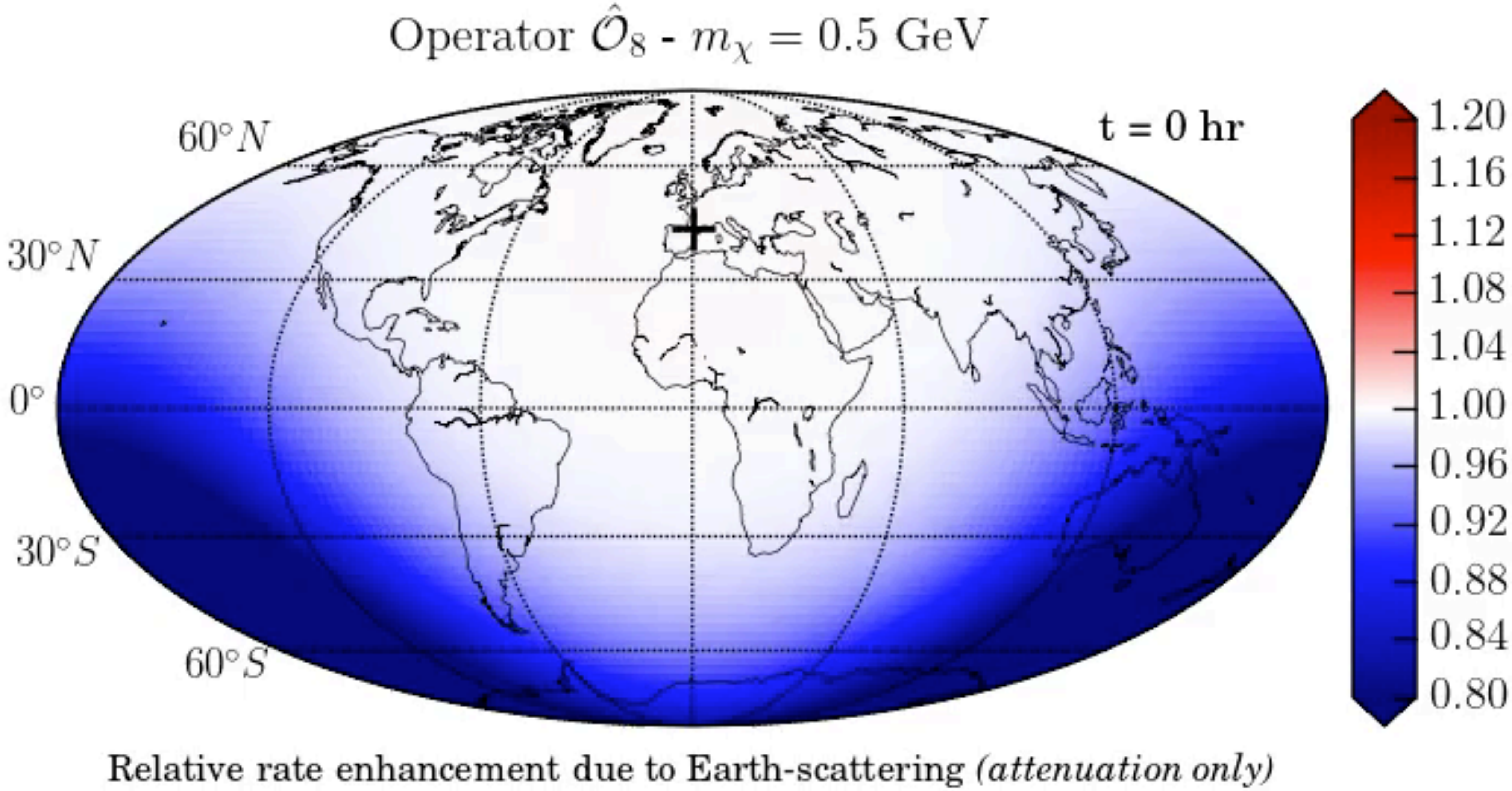


Relative rate enhancement due to Earth-scattering *(attenuation + deflection)*

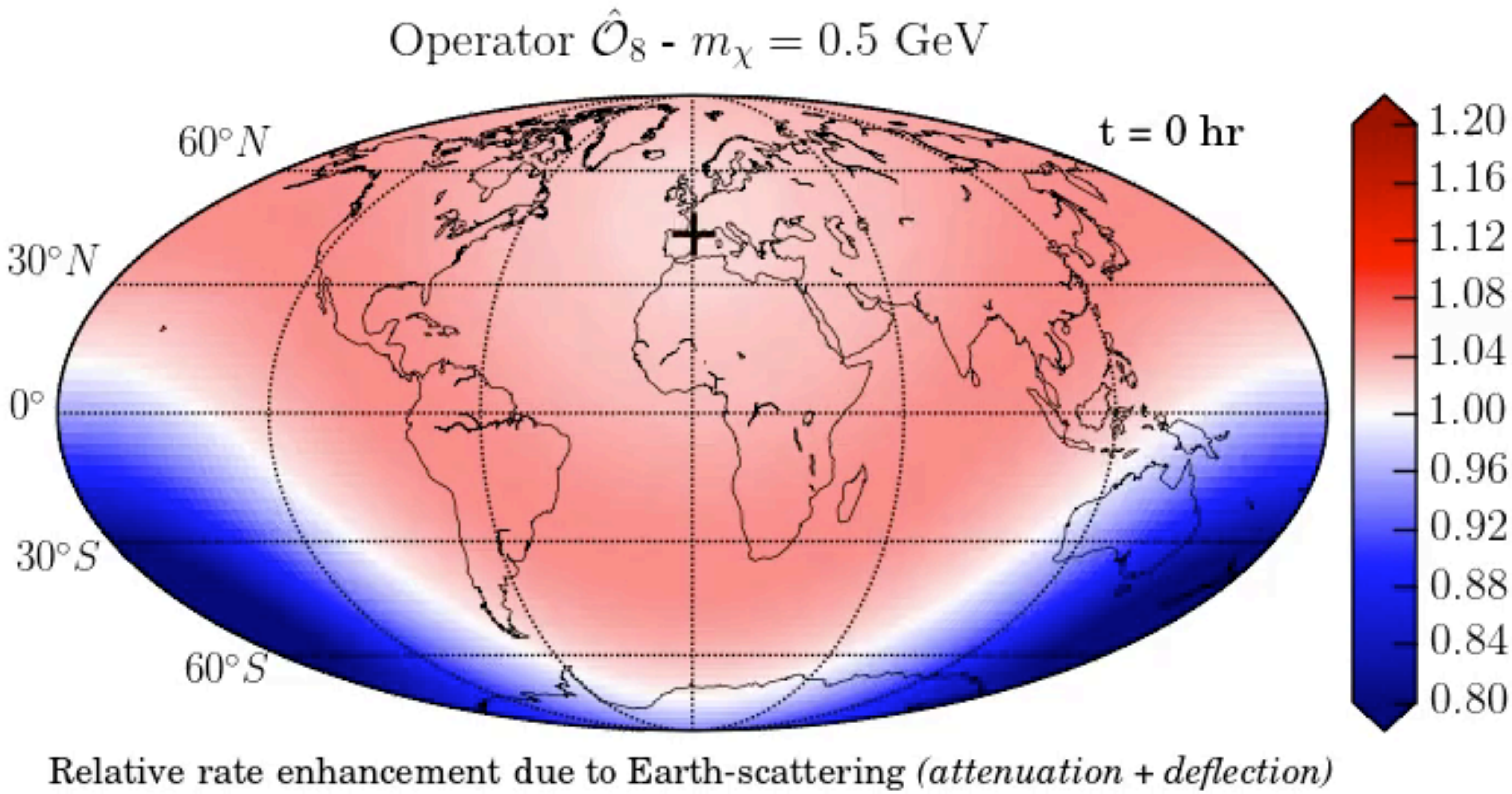
Operator O8



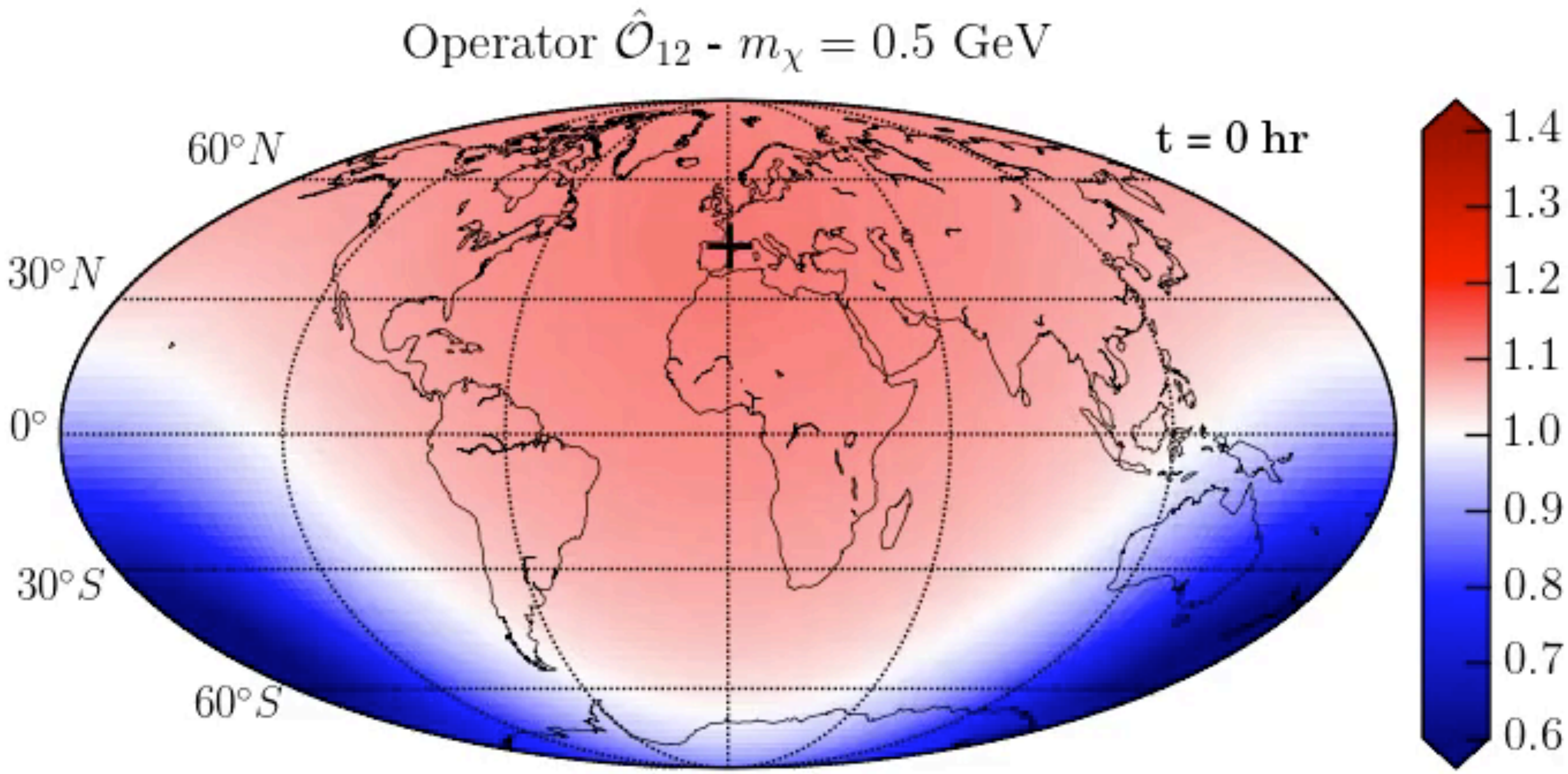
Mapping the CRESST-II Rate



Mapping the CRESST-II Rate



Mapping the CRESST-II Rate



Relative rate enhancement due to Earth-scattering ***(attenuation + deflection)***