

Quantum scale symmetry  
in  
particle physics, gravity  
and cosmology

# Quantum scale symmetry

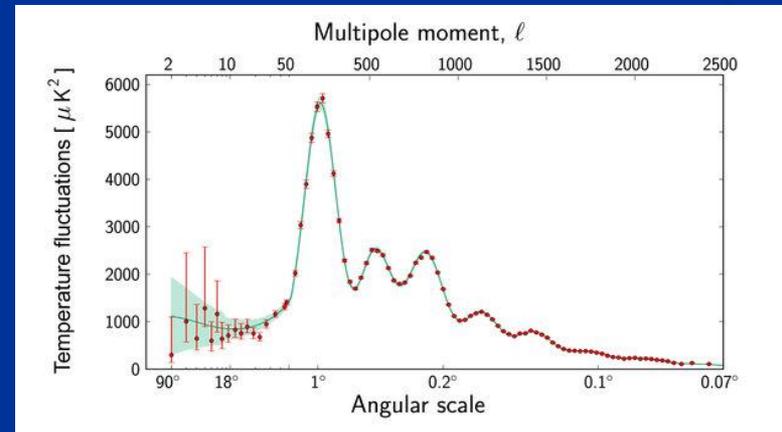
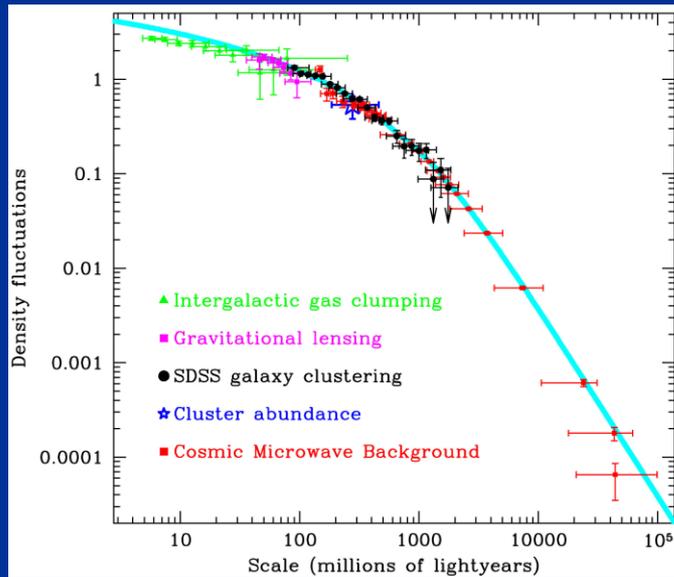
No parameter with dimension of length or mass is present in the quantum effective action.

Then invariance under dilatations or global scale transformations is realized.

Continuous global symmetry

# Scale symmetry in cosmology ?

*Almost scale invariant primordial fluctuation spectrum*



scales are present in cosmology

# Scale symmetry in elementary particle physics ?

proton mass , electron mass

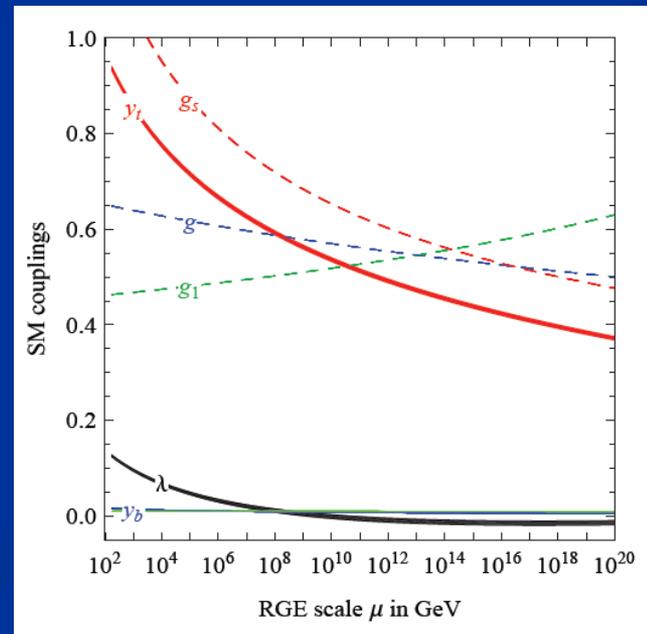
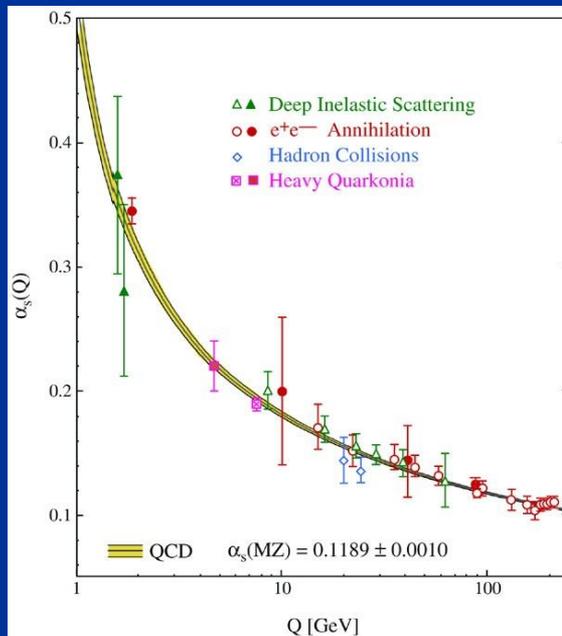
Scales are present in particle physics, but very small as compared to Planck mass

High momentum scattering almost scale invariant

# Quantum scale symmetry

# Quantum fluctuations induce running couplings

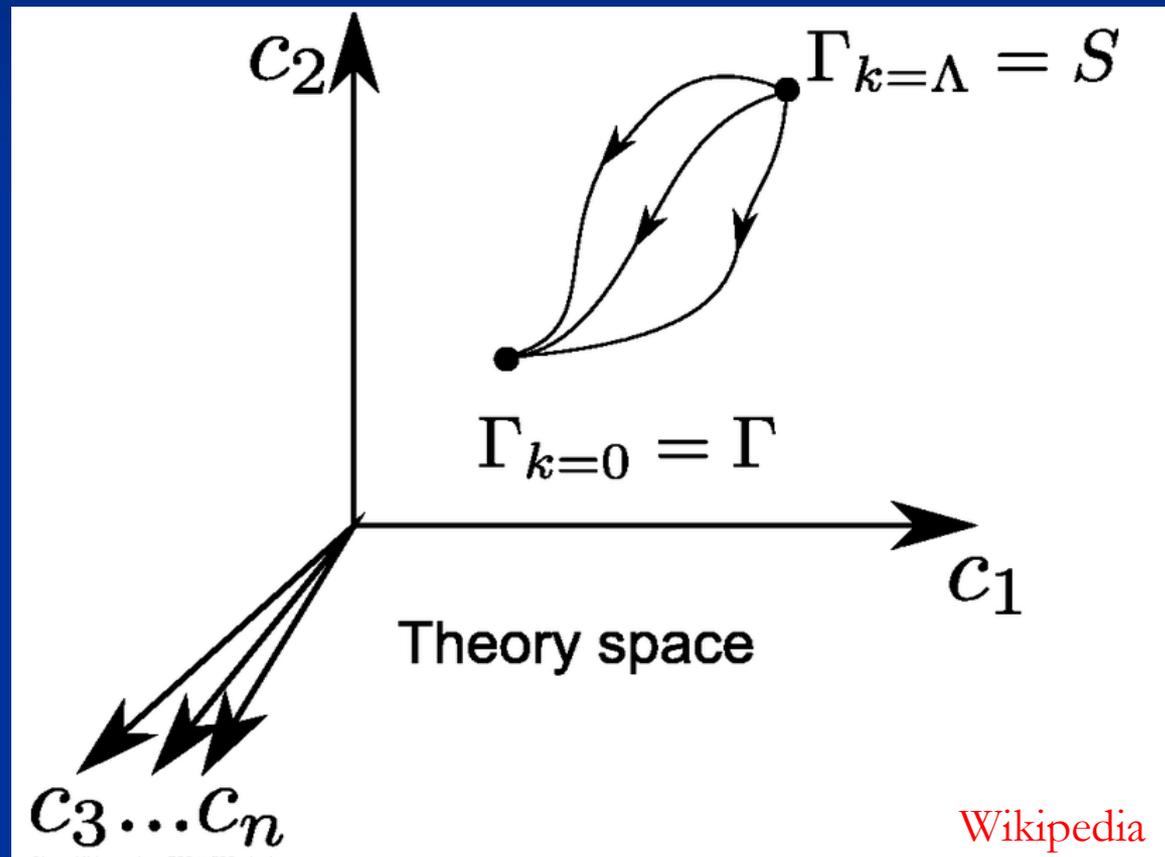
- violation of scale symmetry
- well known in QCD or standard model



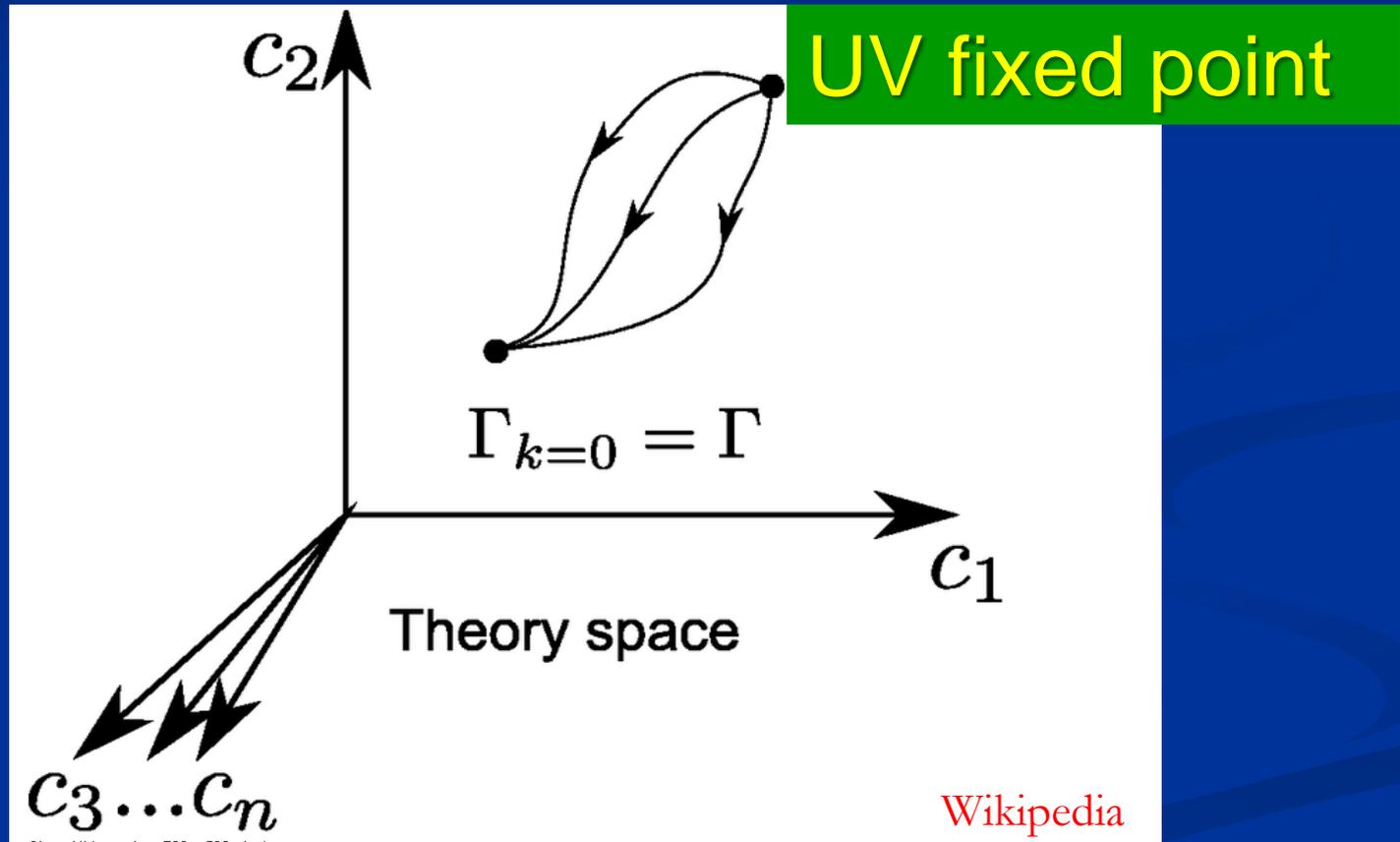
# Quantum scale symmetry

- quantum fluctuations violate scale symmetry
- running dimensionless couplings
- at fixed points , scale symmetry is exact !
- quantum fluctuations can generate scale symmetry !

# Functional renormalization : flowing action



# Ultraviolet fixed point



# Quantum scale symmetry

Exactly on fixed point:

No parameter with dimension of length or mass is present in the quantum effective action.

Then invariance under dilatations or global scale transformations is realized as a quantum symmetry.

Continuous global symmetry

# Two scale symmetries

- Gravity scale symmetry:

includes transformation of metric and scalar singlet

- Particle scale symmetry:

metric and scalar singlet kept fixed

**relative** scaling of momenta with respect to

Planck mass

# *Gravity scale symmetry*



# Gravity scale symmetry

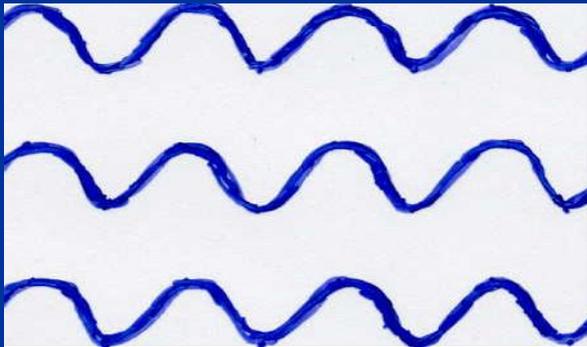
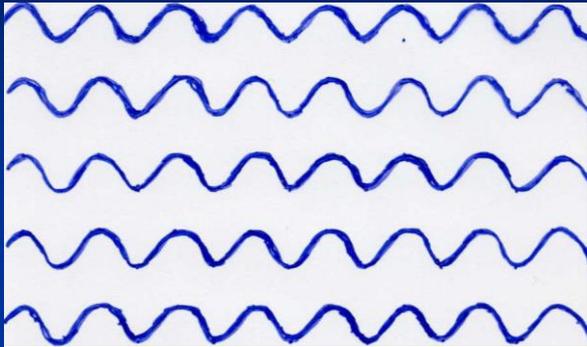
*Replace Planck mass by scalar field*

$$\Gamma = \int_x \sqrt{g} \left\{ -\frac{1}{2} \chi^2 R + \dots \right\}$$

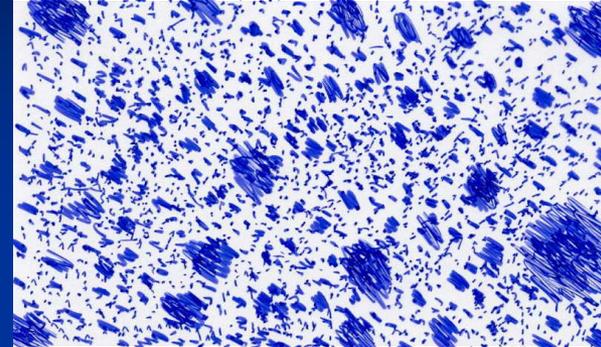
# Spontaneous breaking of scale symmetry

- expectation value of scalar field breaks scale symmetry spontaneously
- massive particles are compatible with scale symmetry
- in presence of massive particles : sign of exact scale symmetry is exactly **massless Goldstone boson** – the dilaton

# Scale symmetry



no scale symmetry



scale symmetry

*only if no spontaneous symmetry breaking!*

# Scale symmetric standard model

## ■ Replace all mass scales by scalar field $\chi$

(1) Higgs potential

$$U = \frac{\lambda_H}{2} (\varphi^\dagger \varphi - \epsilon \chi^2)^2$$



$$\varphi_0^2 = \epsilon \chi^2$$

Fujii,  
Zee, CW

(2) Strong gauge coupling, normalized at  $\mu = \chi$ , is independent of  $\chi$

$$g(\chi) = \bar{g}$$



$$\Lambda_{\text{QCD}} = \chi \exp\left(-\frac{1}{b_0 \bar{g}^2}\right)$$

$$b_0 = \frac{1}{16\pi^2} \left(22 - \frac{4}{3} N_f\right)$$

(3) Similar for all dimensionless couplings

*Quantum effective action for standard model does  
not involve intrinsic mass or length*

Quantum scale symmetry

CW'87

For  $\chi_0 \neq 0$  massless Goldstone boson

# Gravity scale symmetry does not protect small Fermi scale

Effective potential

$$U = \frac{\lambda_H}{2}(\varphi^\dagger\varphi - \epsilon\chi^2)^2$$

is scale invariant for arbitrary  $\epsilon$

$$\varphi_0^2 = \epsilon\chi^2$$

# *Particle scale symmetry*

*Second order  
vacuum electroweak phase transition*



*fixed point*



*quantum scale symmetry*

# Scale symmetry and Fermi scale

- Vacuum electroweak phase transition is (almost) second order, **including all effects from quantum fluctuations**
- Critical surface of second order phase transition: exact fixed point, **quantum scale symmetry**
- Scale symmetry guarantees **“naturalness”** of gauge hierarchy

C. Wetterich, Phys. Lett. B140(1984)215,  
W. A. Bardeen, FERMILAB-CONF-95-391-T(1995)

# Scale symmetry and Fermi scale

- Vacuum electroweak phase transition is (almost) second order
- Critical surface of second order phase transition: exact fixed point, **quantum scale symmetry**
- Scale symmetry guarantees “**naturalness**” of small  $\varphi_0/M$  (gauge hierarchy)

C. Wetterich, Phys. Lett.B140(1984)215,  
W. A. Bardeen, FERMILAB-CONF-95-391-T(1995)

- No fine tuning for renormalisation group improved perturbation theory for deviation from critical surface

$$\mu \frac{\partial}{\partial \mu} \delta = A \delta$$

$$A = \frac{1}{16\pi^2} \left( 2\lambda_H + 6h_t^2 - \frac{9}{2}g_2^2 - \frac{3}{2}g_1^2 \right)$$

# Fine tuning ?

Fine tuning of parameters , quadratic divergences concern bare perturbation theory for location of critical surface in coupling constant space.

not relevant for observation,

not particularly interesting,

regularization dependent, not universal,

always depends on unknown microscopic details

bare perturbation theory is bad expansion

# Quantum Gravity

*Quantum Gravity is a  
renormalisable quantum field theory*

Asymptotic safety

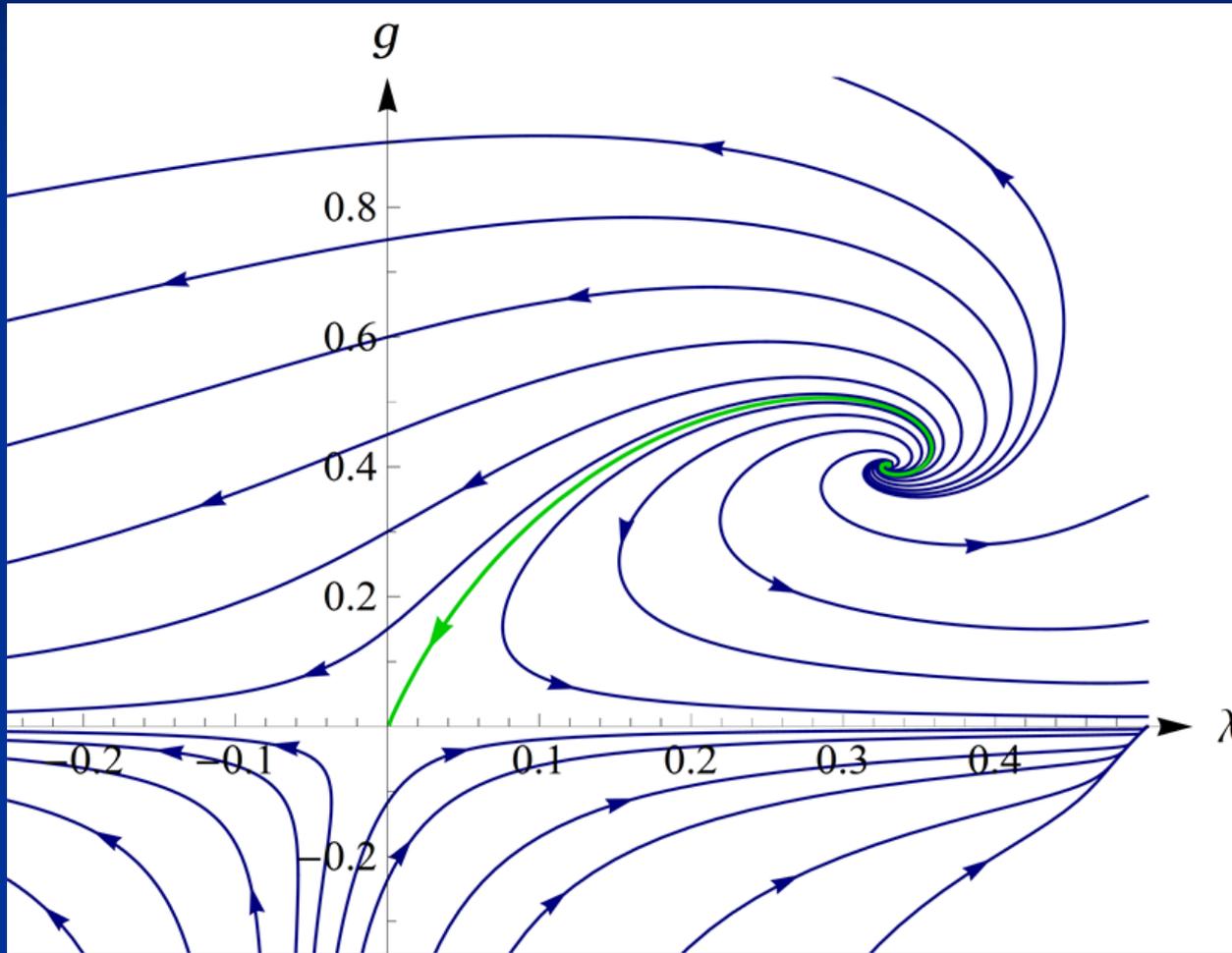
# Asymptotic safety of quantum gravity

if UV fixed point exists :

*quantum gravity is  
non-perturbatively renormalizable !*

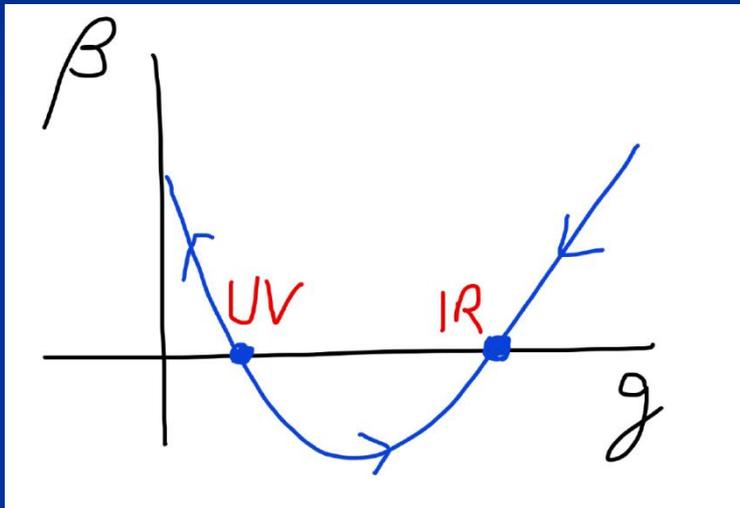
S. Weinberg , M. Reuter

# UV- fixed point for quantum gravity

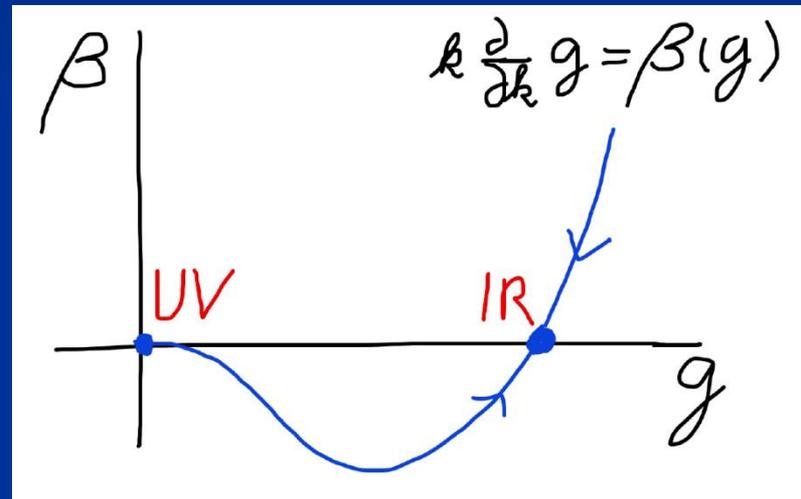


Wikipedia

## Asymptotic safety



## Asymptotic freedom



Relevant parameters yield undetermined couplings. Quartic scalar coupling is not relevant and can therefore be predicted.

# a prediction...

## Asymptotic safety of gravity and the Higgs boson mass

Mikhail Shaposhnikov

*Institut de Théorie des Phénomènes Physiques, École Polytechnique Fédérale de Lausanne, CH-1015 Lausanne, Switzerland*

Christof Wetterich

*Institut für Theoretische Physik, Universität Heidelberg, Philosophenweg 16, D-69120 Heidelberg, Germany*

12 January 2010

### Abstract

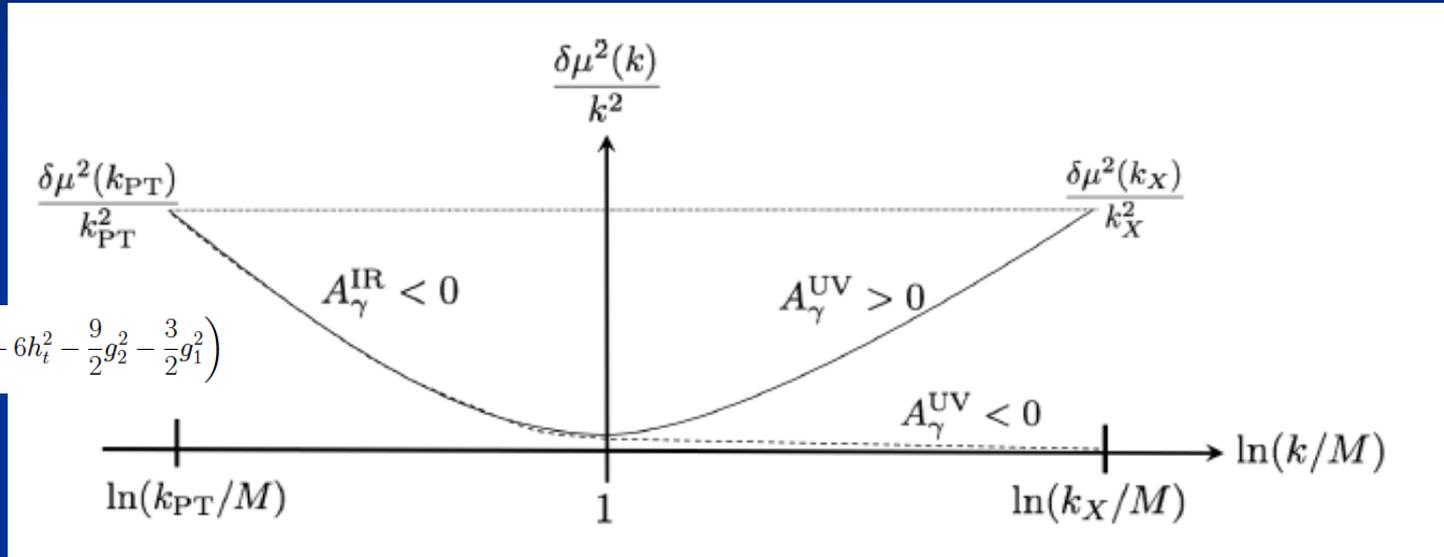
There are indications that gravity is asymptotically safe. The Standard Model (SM) plus gravity could be valid up to arbitrarily high energies. Supposing that this is indeed the case and assuming that there are no intermediate energy scales between the Fermi and Planck scales we address the question of whether the mass of the Higgs boson  $m_H$  can be predicted. For a positive gravity induced anomalous dimension  $A_\lambda > 0$  the running of the quartic scalar self interaction  $\lambda$  at scales beyond the Planck mass is determined by a fixed point at zero. This results in  $m_H = m_{\min} = 126$  GeV, with only a few GeV uncertainty. This prediction is independent of the details of the short distance running and holds for a wide class of extensions of the SM as well.

s in  $m_H = m_{\min} = 126$  GeV, with o

# Possible explanation of gauge hierarchy

$$A_\gamma^{\text{IR}} = -2 + A$$

$$A = \frac{1}{16\pi^2} \left( 2\lambda_H + 6h_t^2 - \frac{9}{2}g_2^2 - \frac{3}{2}g_1^2 \right)$$



Gauge hierarchy problem in asymptotically safe gravity  
–the resurgence mechanism

Christof Wetterich<sup>1</sup> and Masatoshi Yamada<sup>1</sup>

Phys.Lett. B770 (2017) 268-271

Quantum gravity with  
scalar field –  
the role of scale symmetry  
for cosmology

# Exact scale symmetry ?

*Precisely on fixed point :*

- Exact scale symmetry

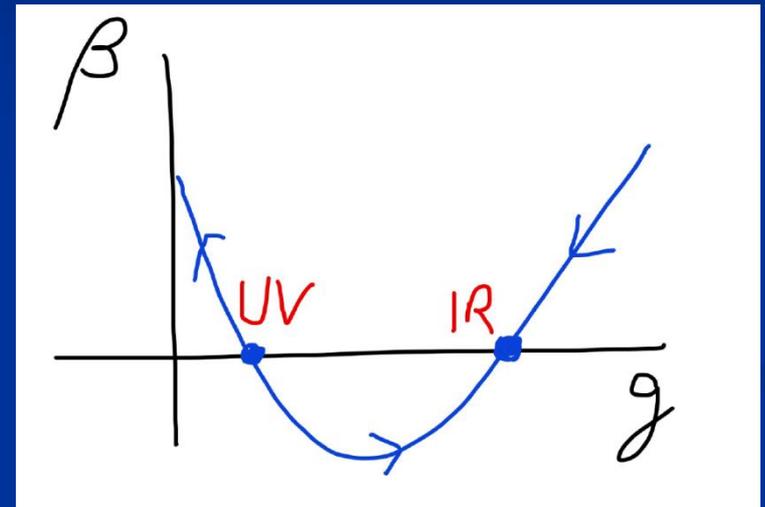
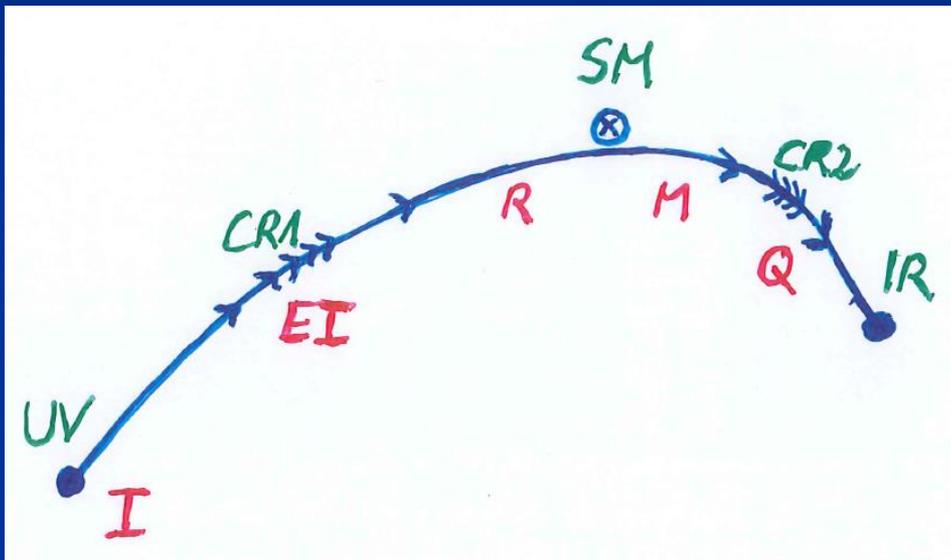
*Vicinity of fixed point :*

- Relevant parameters induce intrinsic scales by flow away from fixed point.
- Approximate scale symmetry in vicinity of fixed point.
- Fixed point with exact scale symmetry only reached in extreme UV or IR limit.

# Approximate scale symmetry near fixed points

- UV : approximate scale invariance of primordial fluctuation spectrum from inflation
- IR : cosmon is pseudo Goldstone boson of spontaneously broken scale symmetry, tiny mass, responsible for dynamical Dark Energy

# Possible consequences of crossover in quantum gravity



Realistic model for inflation and dark energy  
with single scalar field

# variable gravity

*“Newton’s constant is not constant –  
and particle masses are not constant”*

Πάντα ρεῖ



# Variable Gravity

$$\Gamma = \int_x \sqrt{g} \left\{ -\frac{1}{2} \chi^2 R + \mu^2 \chi^2 + \frac{1}{2} (B(\chi/\mu) - 6) \partial^\mu \chi \partial_\mu \chi \right\}$$

quantum effective action,  
variation yields field equations

Einstein gravity :  $\Gamma = \int_x \sqrt{g} \left\{ -\frac{1}{2} M^2 R \right\}$

# + scale symmetric standard model

- Replace all mass scales by scalar field  $\chi$

(1) Higgs potential

$$U = \frac{\lambda_H}{2}(\varphi^\dagger\varphi - \epsilon\chi^2)^2 \quad \longrightarrow \quad \varphi_0^2 = \epsilon\chi^2$$

(2) Strong gauge coupling, normalized at  $\mu = \chi$ , is independent of  $\chi$

$$g(\chi) = \bar{g} \quad \longrightarrow \quad \Lambda_{\text{QCD}} = \chi \exp\left(-\frac{1}{b_0\bar{g}^2}\right) \quad b_0 = \frac{1}{16\pi^2} \left(22 - \frac{4}{3}N_f\right)$$

# + scale invariant action for dark matter

# Scale symmetry in variable gravity ( IR – fixed point )

$$\Gamma = \int_x \sqrt{g} \left\{ -\frac{1}{2} \chi^2 R + \cancel{\mu^2 \chi^2} + \frac{1}{2} (\cancel{B(\chi/\mu)} - 6) \partial^\mu \chi \partial_\mu \chi \right\}$$

quantum effective action,  
variation yields field equations

Einstein gravity :  $\Gamma = \int_x \sqrt{g} \left\{ -\frac{1}{2} M^2 R \right\}$

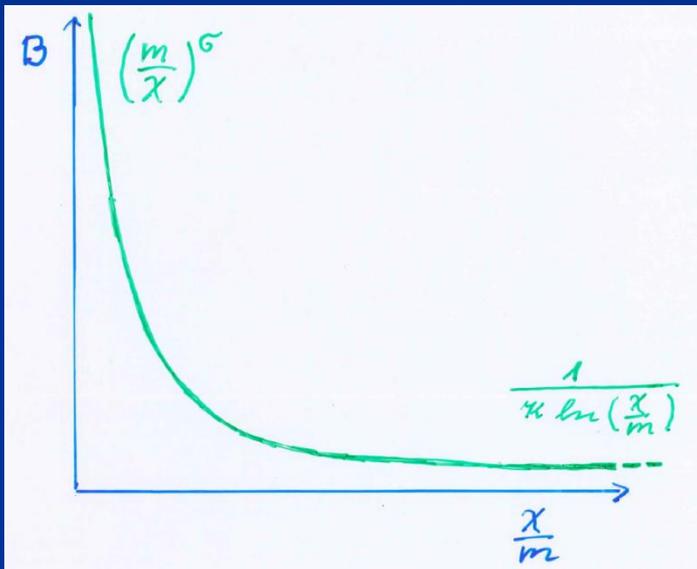
# Variable Gravity

- Scalar field coupled to gravity
- Effective Planck mass depends on scalar field
- Simple quadratic scalar potential involves intrinsic mass  $\mu$
- Nucleon and electron mass proportional to dynamical Planck mass

$$\Gamma = \int_x \sqrt{g} \left\{ -\frac{1}{2}\chi^2 R + \mu^2\chi^2 + \frac{1}{2}(B(\chi/\mu) - 6)\partial^\mu\chi\partial_\mu\chi \right\}$$

# Kinetic B :

## Crossover between two fixed points



assumption:

running

coupling obeys

flow equation

$$\mu \frac{\partial B}{\partial \mu} = \frac{\kappa \sigma B^2}{\sigma + \kappa B}$$

$$B^{-1} - \frac{\kappa}{\sigma} \ln B = \kappa \left[ \ln \left( \frac{\chi}{\mu} \right) - c_t \right] = \kappa \ln \left( \frac{\chi}{m} \right)$$

$m$  : scale of crossover

can be exponentially larger than intrinsic scale  $\mu$

# Four-parameter model

- model has four dimensionless parameters
- three in kinetial :
  - $\sigma \sim 2.5$
  - $\kappa \sim 0.5$
  - $c_t \sim 14$  ( or  $m/\mu$  )
- one parameter for growth rate of neutrino mass over electron mass :  $\gamma \sim 8$
- + standard model particles and dark matter : sufficient for realistic cosmology from inflation to dark energy
- no more free parameters than  $\Lambda$ CDM

# Cosmology

*Add matter and radiation*

*(standard model + dark matter)*

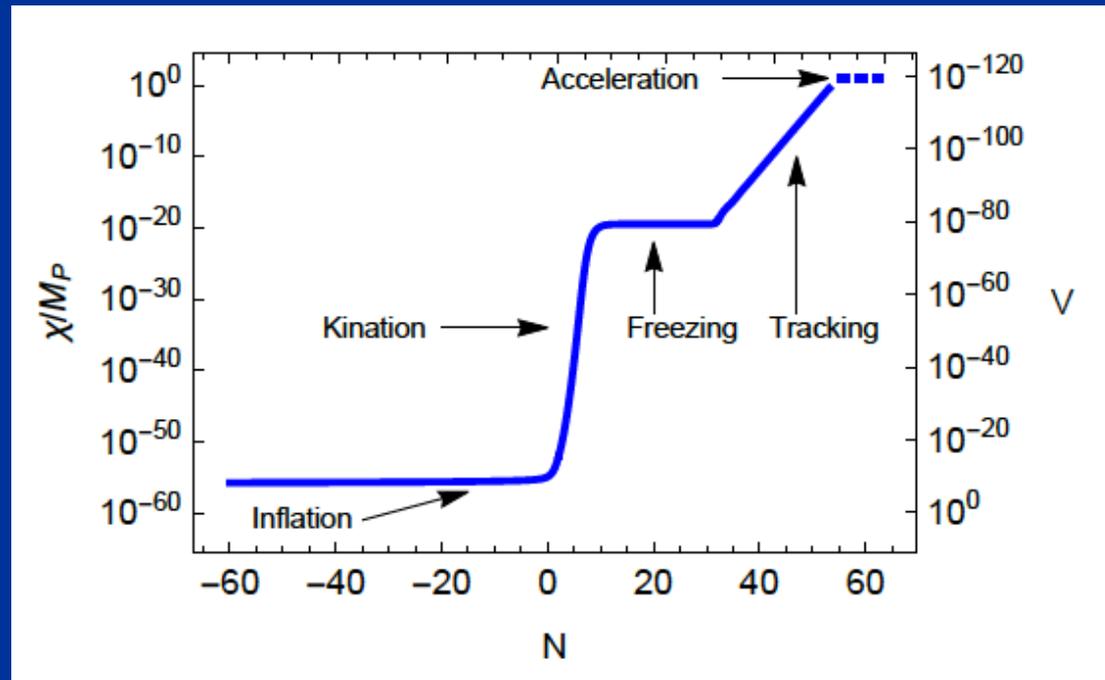
*Solve field equations...*

$$\Gamma = \int_x \sqrt{g} \left\{ -\frac{1}{2} \chi^2 R + \mu^2 \chi^2 + \frac{1}{2} (B(\chi/\mu) - 6) \partial^\mu \chi \partial_\mu \chi \right\}$$

$$B^{-1} - \frac{\kappa}{\sigma} \ln B = \kappa \left[ \ln \left( \frac{\chi}{\mu} \right) - c_t \right] = \kappa \ln \left( \frac{\chi}{m} \right)$$

# Cosmological solution

- scalar field  $\chi$  vanishes in the infinite past
- scalar field  $\chi$  diverges in the infinite future



# Model is compatible with present observations

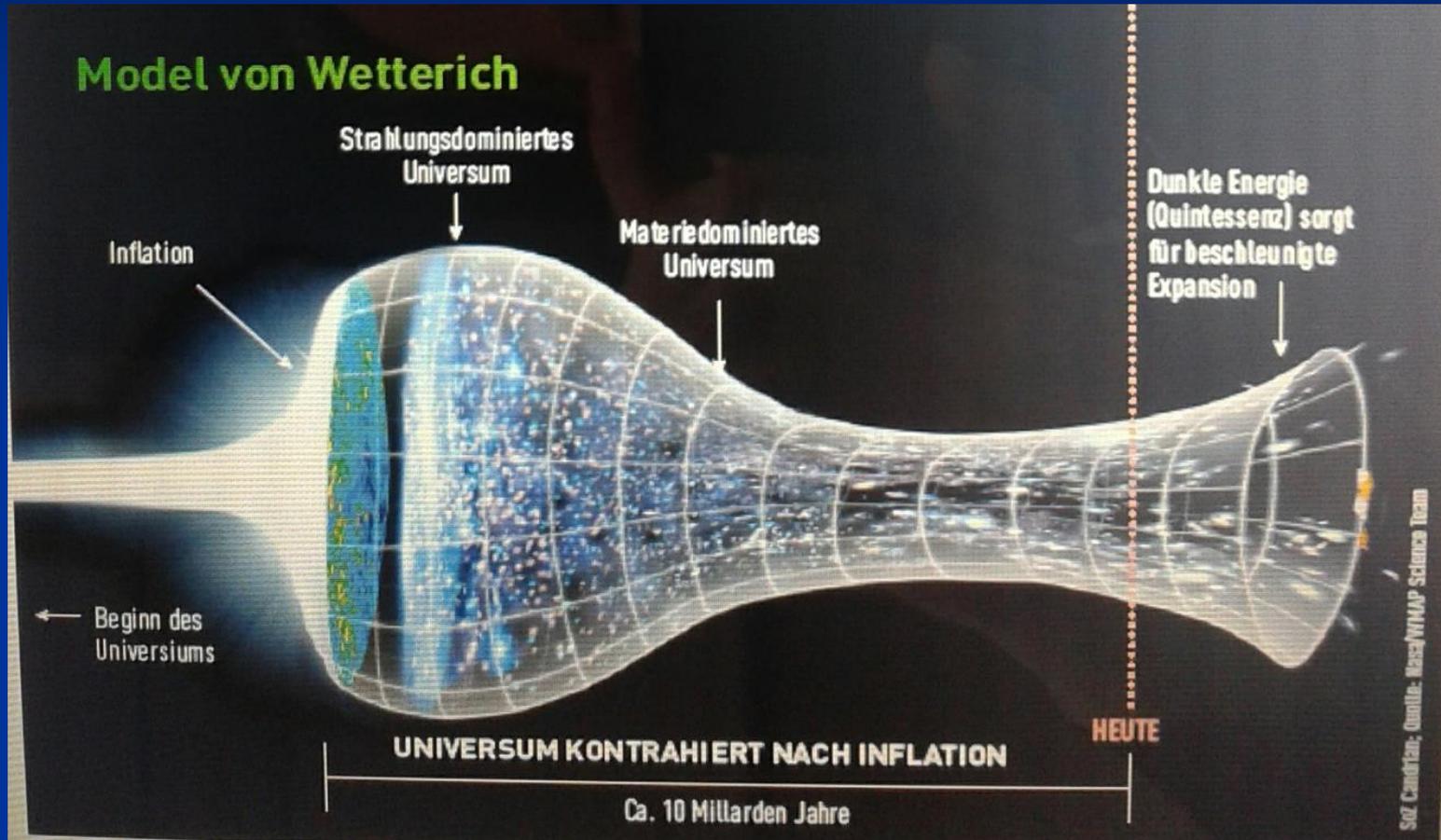
Together with variation of neutrino mass over electron mass in present cosmological epoch :

model is compatible with all present observations, including inflation and dark energy

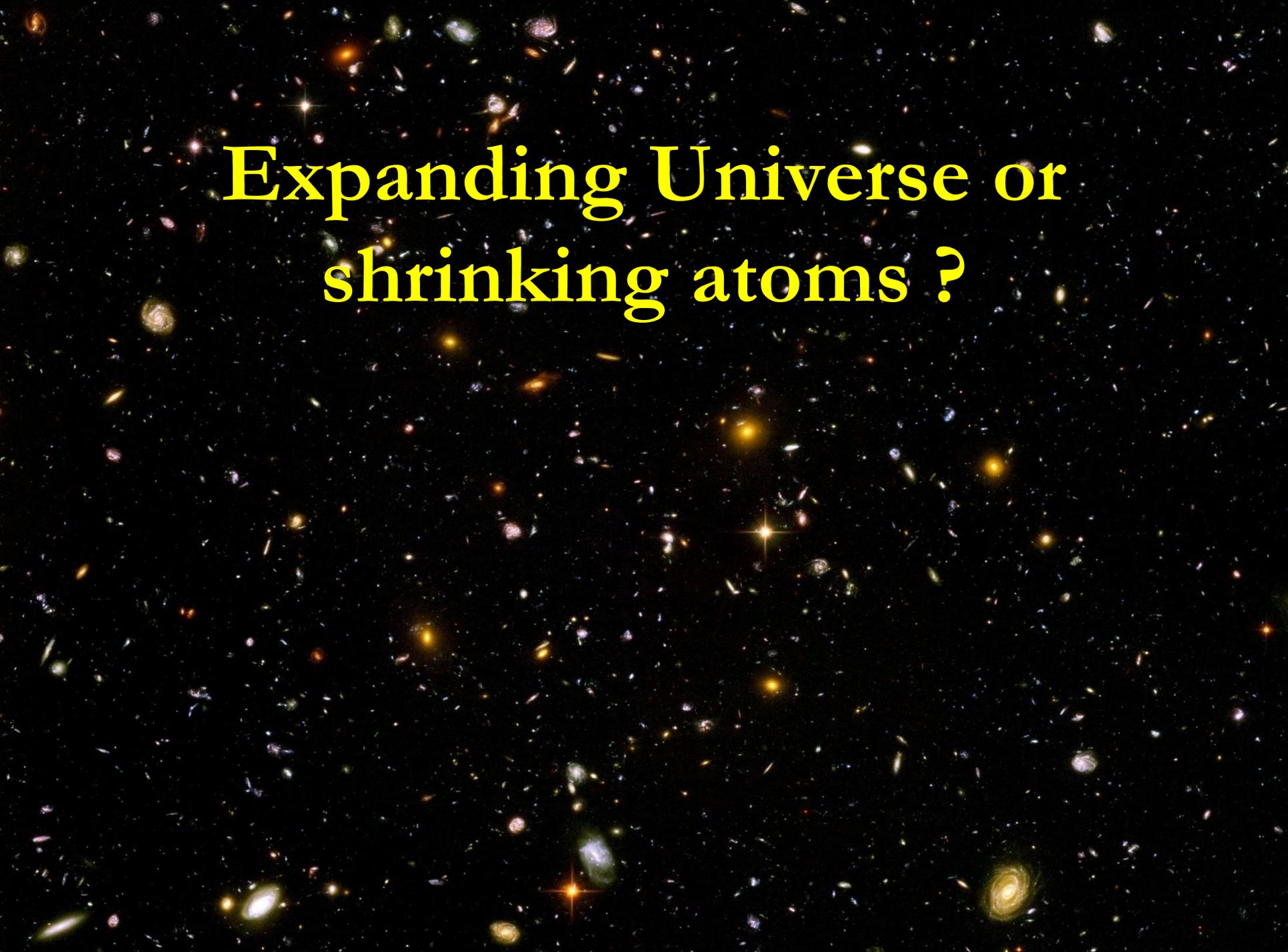
$$\Gamma = \int_x \sqrt{g} \left\{ -\frac{1}{2}\chi^2 R + \mu^2\chi^2 + \frac{1}{2}(B(\chi/\mu) - 6)\partial^\mu\chi\partial_\mu\chi \right\}$$

$$B^{-1} - \frac{\kappa}{\sigma} \ln B = \kappa \left[ \ln \left( \frac{\chi}{\mu} \right) - c_t \right] = \kappa \ln \left( \frac{\chi}{m} \right)$$

# Strange evolution of Universe



Sonntagszeitung Zürich , Laukenmann

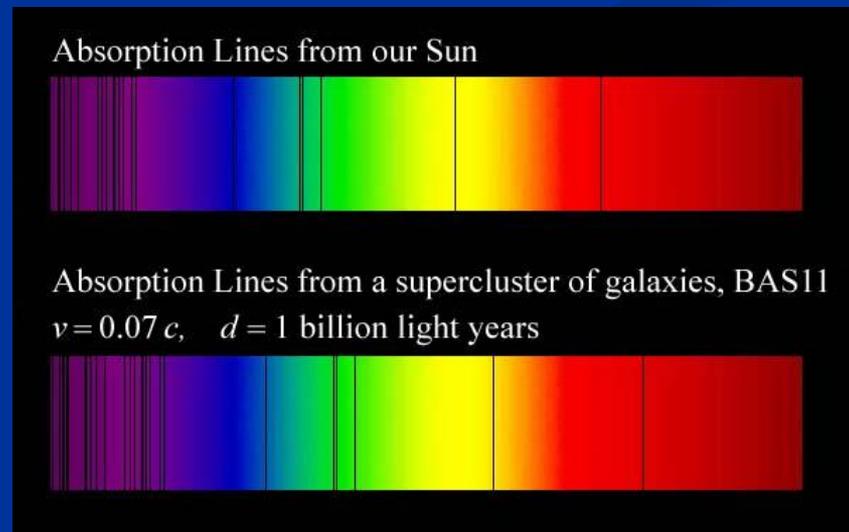
The background of the image is a vast field of galaxies, likely from a deep space survey. The galaxies are scattered across the frame, appearing in various colors including yellow, orange, red, blue, and purple. Some are bright and clear, while others are faint and distant. The overall effect is a rich, multi-colored cosmic landscape.

**Expanding Universe or  
shrinking atoms ?**

# Do we know that the Universe expands ?

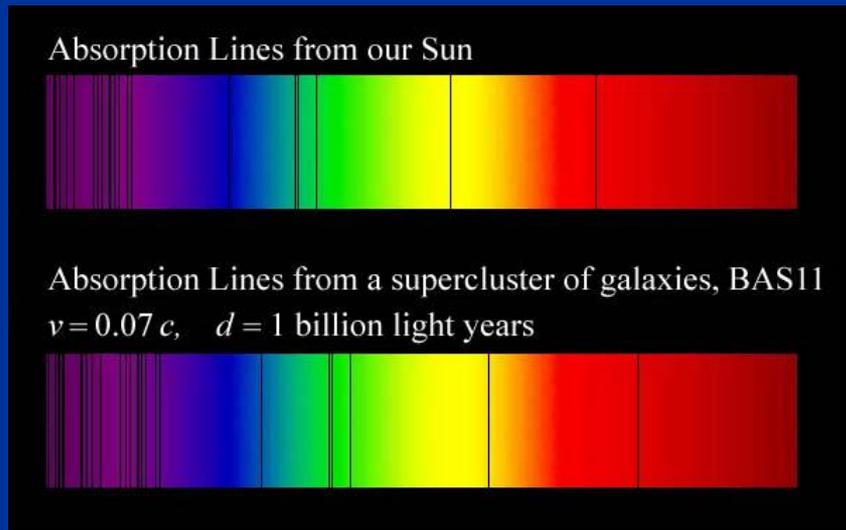
instead of redshift due to expansion :

smaller frequencies have been emitted in the past,  
because electron mass was smaller !



# Why do we see redshift of photons emitted in the distant past ?

photons are more red because they have been **emitted** with longer wavelength



frequency  $\sim$  mass

wavelength  $\sim$   
atoms size

# What is increasing ?

Ratio of distance between galaxies  
over size of atoms !

atom size constant : expanding geometry

alternative : shrinking size of atoms

# How can particle masses change with time ?

- Particle masses are proportional to scalar field  $\chi$  .  
Similar to Higgs field.
- Scalar field varies with time.
- Ratios of particle masses are independent of  $\chi$  and therefore remain constant.
- Compatibility with observational bounds on time dependence of particle mass ratios.
- Dimensionless couplings are independent of  $\chi$  .

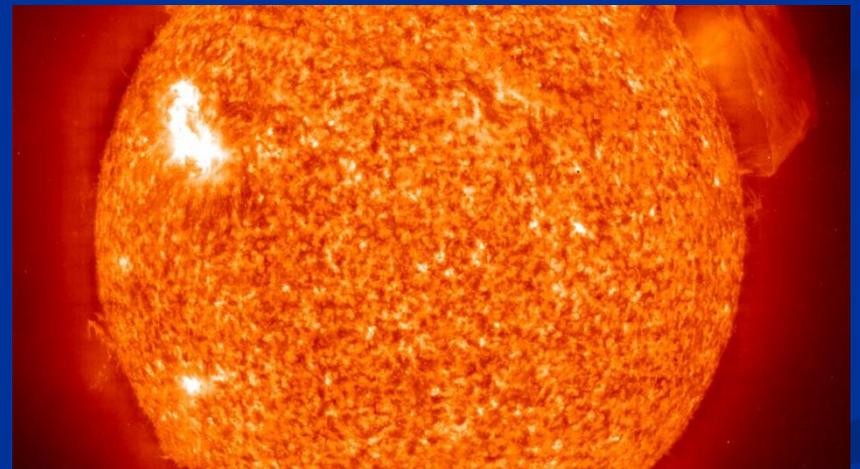
A vast field of galaxies, including spirals, ellipticals, and irregular shapes, scattered across a dark background. The galaxies are in various colors, including yellow, orange, blue, and purple. The text "Hot big bang or freeze?" is overlaid in the center in a yellow, serif font.

**Hot big bang or freeze ?**

Do we know that the temperature was higher in the early Universe than now ?

Cosmic microwave radiation , nucleosynthesis

instead of  
higher temperature :  
smaller particle masses



# Hot plasma ?

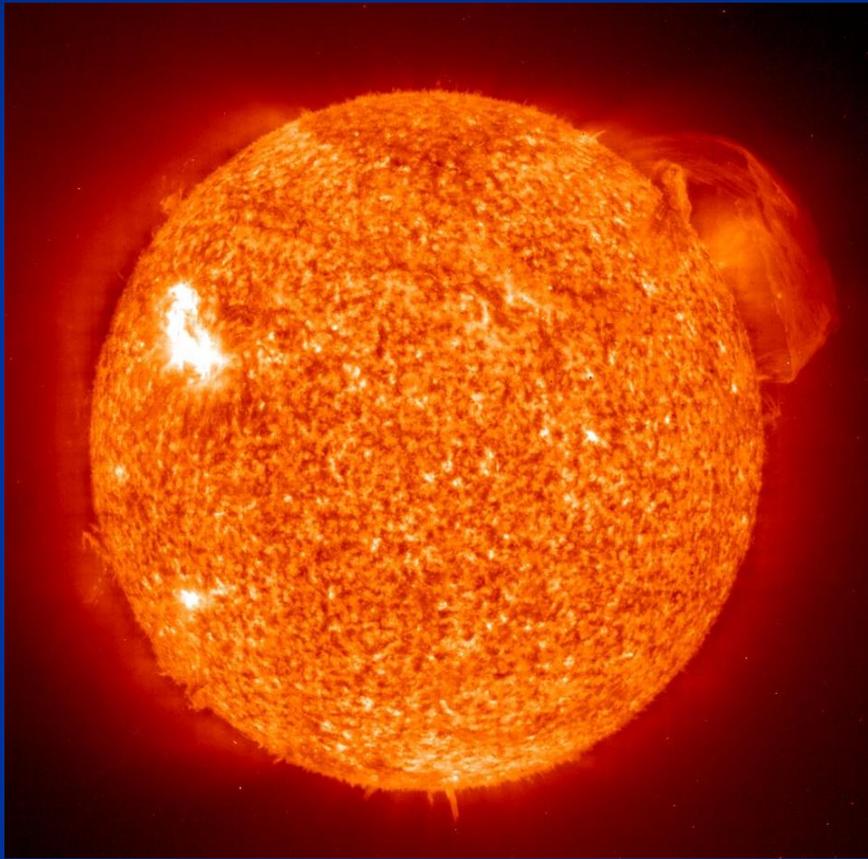
- Temperature in radiation dominated Universe :  
 $T \sim \chi^{1/2}$  **smaller** than today
- Ratio temperature / particle mass :  
 $T / m_p \sim \chi^{-1/2}$  **larger** than today
- $T/m_p$  counts ! This ratio decreases with time.
- **Nucleosynthesis , CMB emission as in standard cosmology !**

# Freeze Universe

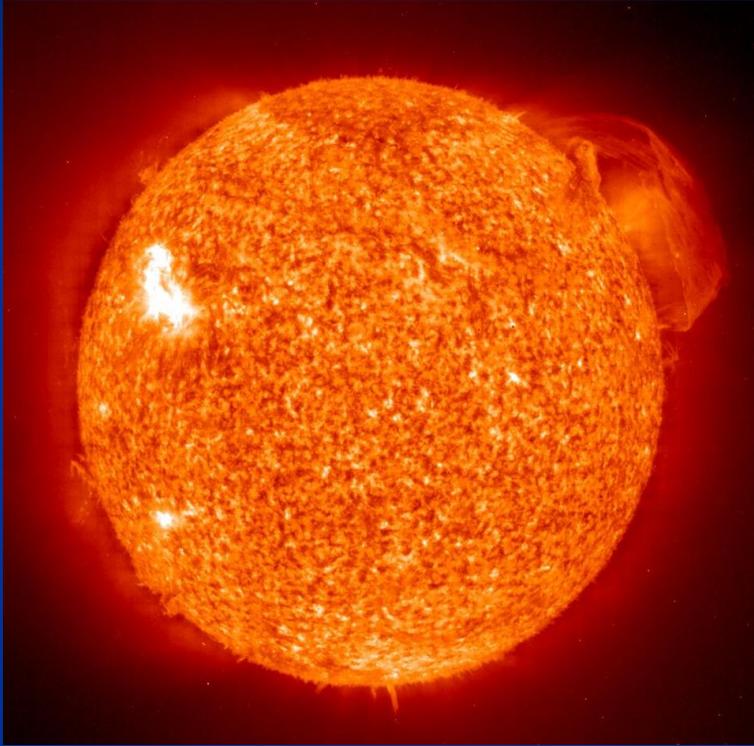
*The Universe may have started very cold,  
and only later heat up.*

*Freeze picture of the Universe*

# Big bang or freeze ?



# Big bang or freeze ?



freeze picture :  
only rods for measurements  
( masses ) are different !

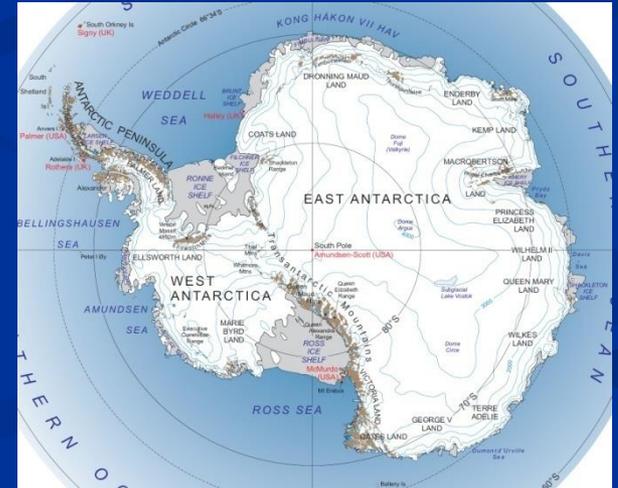
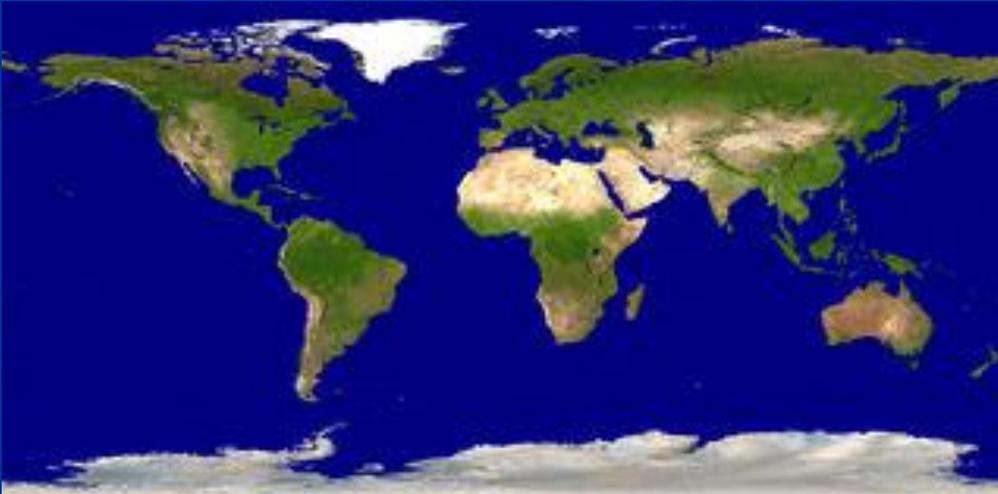
*Big bang is not wrong,*

*but alternative pictures exist !*

# Field - singularity

- Big Bang is field - singularity
- similar ( but not identical with )  
coordinate - singularity

$$g'_{\mu\nu} = \frac{\chi^2}{M^2} g_{\mu\nu}$$



# Conclusions

Quantum gravity may be observable in  
dynamics of present Universe

Fixed points and scale symmetry crucial

Big bang singularity is artefact  
of inappropriate choice of field variables –  
no physical singularity



end

# Einstein frame

- “Weyl scaling” maps variable gravity model to Universe with fixed masses and standard expansion history.
- Exact equivalence of different frames !  
“different pictures”
- Standard gravity coupled to scalar field.

# Einstein frame

Weyl scaling :

$$g'_{\mu\nu} = \frac{\chi^2}{M^2} g_{\mu\nu} , \quad \varphi = \frac{2M}{\alpha} \ln \left( \frac{\chi}{\mu} \right)$$

effective action in Einstein frame :

$$\Gamma = \int_x \sqrt{g'} \left\{ -\frac{1}{2} M^2 R' + V'(\varphi) + \frac{1}{2} k^2(\varphi) \partial^\mu \varphi \partial_\mu \varphi \right\}$$

$$V'(\varphi) = M^4 \exp \left( -\frac{\alpha \varphi}{M} \right)$$

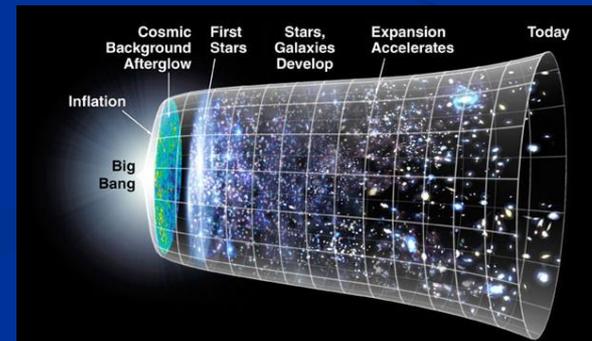
$$k^2 = \frac{\alpha^2 B}{4}$$

# Field relativity

Weyl scaling :

$$g'_{\mu\nu} = \frac{\chi^2}{M^2} g_{\mu\nu}$$

changes geometry,  
not a coordinate transformation



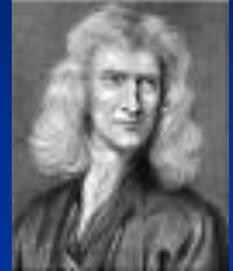
# Field relativity :

## different pictures of cosmology

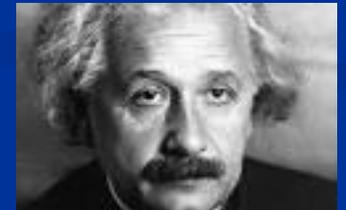
- same physical content can be described by different pictures
- related by field – redefinitions ,  
e.g. Weyl scaling , conformal scaling of metric
- observables cannot depend on choice of fields
- metric is one of the fields
- which picture is usefull ?

# Relativity of geometry

- **Euclid ... Newton** : space and time are absolute



- **Special relativity** : space and time depend on observer
- **General relativity** : space-time is influenced by matter ( including radiation )  
geometry is independent of coordinates  
geometry is observable
- **Field relativity** : geometry is relative



*Space-time is a description  
of correlations between “matter”.*

*Observables cannot depend on choice of  
fields used to describe them.*

*Different pictures for geometry exist.*

# Why should you care about the freeze picture of the Universe ?

*Some aspects are understood easier :*

- Natural tiny present dark energy
- Beginning of Universe
- Role of scale symmetry
- Range of impact of quantum gravity

# What is Dark Energy ?

Dark energy is energy density of scalar field  $\chi$

$$\rho = V + \text{kinetic term}$$

$$V = \mu^2 \chi^2$$

$$p = -V + \text{kinetic term}$$

Dark energy is dynamical if  $\chi$  changes with time

**No small parameter for  
dark energy**

# Four-parameter model

- model has four dimensionless parameters
- three in kineticial :
  - $\sigma \sim 2.5$
  - $\kappa \sim 0.5$
  - $c_t \sim 14$  ( or  $m/\mu$  )
- one parameter for growth rate of neutrino mass over electron mass :  $\gamma \sim 8$
- + standard model particles and dark matter : sufficient for realistic cosmology from inflation to dark energy
- no more free parameters than  $\Lambda$ CDM

# asymptotically vanishing cosmological „constant“

- What matters : Ratio of potential divided by fourth power of Planck mass

$$\frac{V}{\chi^4} = \frac{\mu^2 \chi^2}{\chi^4} = \frac{\mu^2}{\chi^2}$$

$$V = \mu^2 \chi^2$$

- vanishes for  $\chi \rightarrow \infty$  !

# small dimensionless number ?

- needs two intrinsic mass scales
- standard approach :  $V$  and  $M$  ( cosmological constant and Planck mass )
- variable gravity : Planck mass moving to infinity , with fixed or moderately increasing  $V$ 
  - ➔ ratio vanishes asymptotically !

# Quintessence

Dynamical dark energy ,  
generated by scalar field (cosmon )

C.Wetterich,Nucl.Phys.B302(1988)668, 24.9.87

P.J.E.Peebles,B.Ratra,ApJ.Lett.325(1988)L17, 20.10.87

## Prediction :

homogeneous dark energy  
influences recent cosmology

- of same order as dark matter -

Original models do not fit the present observations  
.... modifications  
( different growth of neutrino mass )

# No tiny dimensionless parameters ( except gauge hierarchy )

- one mass scale  $\mu = 2 \cdot 10^{-33} \text{ eV}$
- one time scale  $\mu^{-1} = 10^{10} \text{ yr}$
- Planck mass does not appear as parameter
- Planck mass grows large dynamically
- Dark energy is tiny because Universe is old

# Slow Universe

Asymptotic solution in  
freeze frame :

$$H = \frac{\mu}{\sqrt{3}}, \quad \chi = \frac{3^{\frac{1}{4}} m}{2\sqrt{\mu}} (t_c - t)^{-\frac{1}{2}}$$

$$\mu = 2 \cdot 10^{-33} \text{ eV}$$

Expansion or shrinking always slow ,  
characteristic time scale of the order of the age of the  
Universe :  $t_{\text{ch}} \sim \mu^{-1} \sim 10 \text{ billion years !}$

Hubble parameter of the order of **present** Hubble  
parameter for all times , including inflation and big bang !  
Slow increase of particle masses !

**infinite past**

# Infinite past : slow inflation

$\sigma = 2$  : field equations

$$\ddot{\chi} + \left( 3H + \frac{1}{2} \frac{\dot{\chi}}{\chi} \right) \dot{\chi} = \frac{2\mu^2 \chi^2}{m}$$

$$H = \sqrt{\frac{\mu^2}{3} + \frac{m\dot{\chi}^2}{6\chi^3}} - \frac{\dot{\chi}}{\chi}$$

approximative  
solution

$$H = \frac{\mu}{\sqrt{3}}, \quad \chi = \frac{3^{\frac{1}{4}} m}{2\sqrt{\mu}} (t_c - t)^{-\frac{1}{2}}$$

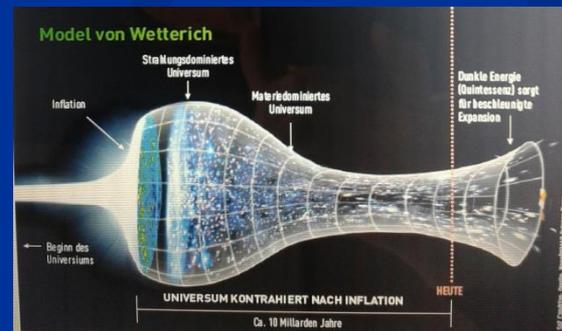
particles become massless in infinite past !

# Eternal Universe

Asymptotic solution in freeze frame :

$$H = \frac{\mu}{\sqrt{3}}, \quad \chi = \frac{3^{\frac{1}{4}} m}{2\sqrt{\mu}} (t_c - t)^{-\frac{1}{2}}$$

- solution valid back to the infinite past in physical time
- no singularity



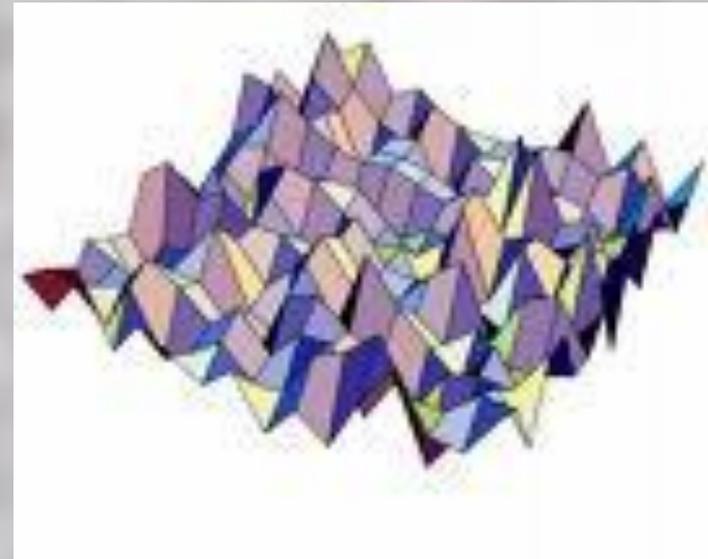
- physical time to infinite past is infinite

# Eternal light-vacuum

Everywhere almost nothing  
only fluctuations

All particles move  
with light velocity,  
similar to photons

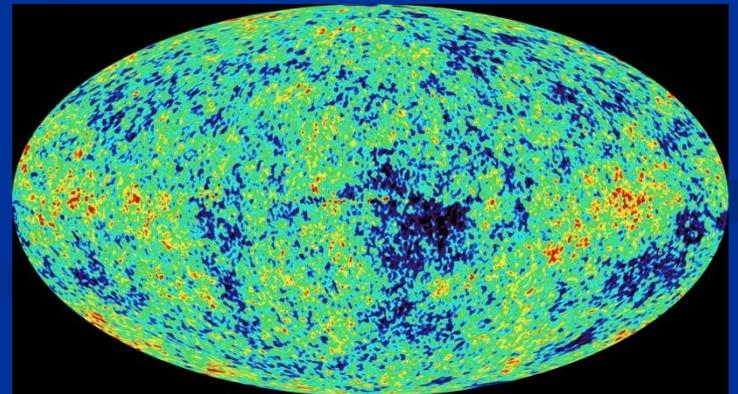
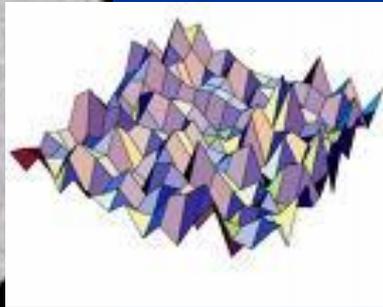
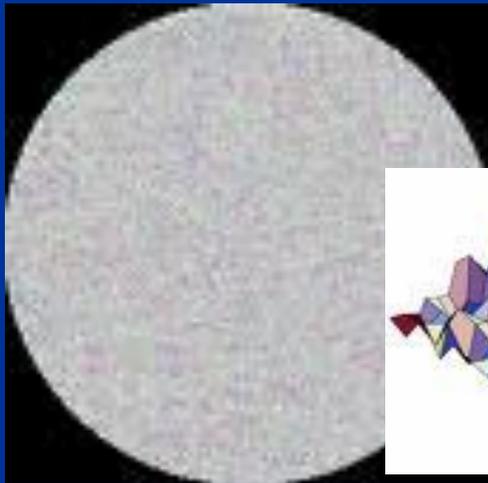
strength of gravity  
much stronger than today



*In the beginning was light-like emptiness.*

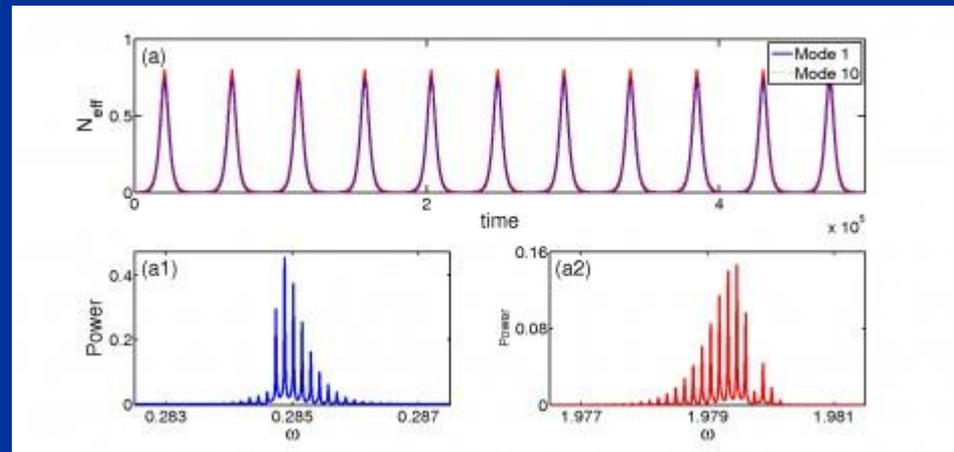
# Eternal light-vacuum is unstable

- Slow increase of particle masses and weakening of gravity
- Only slow change of space-time geometry
- Consequence for observation : primordial fluctuations become visible in cosmic background radiation
- We see fluctuations in a stage 5000 billion years ago.



# Physical time

count oscillations ....



# Physical time

field equation for scalar field mode

$$(\partial_\eta^2 + 2Ha\partial_\eta + k^2 + a^2m^2)\varphi_k = 0$$

$$\varphi_k = \frac{\tilde{\varphi}_k}{a} \quad \left\{ \partial_\eta^2 + k^2 + a^2 \left( m^2 - \frac{R}{6} \right) \right\} \tilde{\varphi}_k = 0$$

determine **physical time** by counting number of oscillations

$$\tilde{t}_p = n_k$$

$$n_k = \frac{k\eta}{\pi}$$

(  $m=0$  )

# Physical time

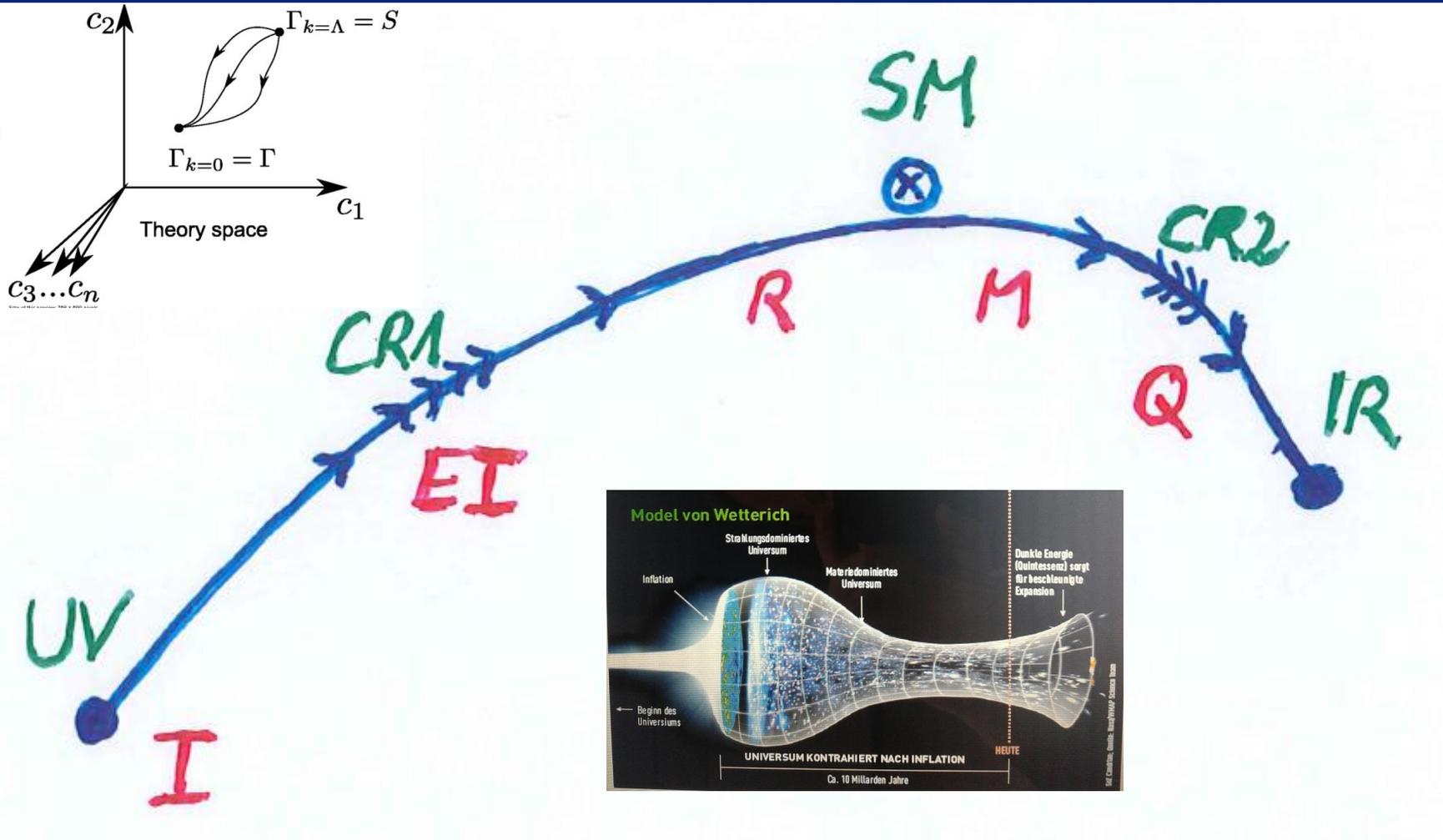
- counting : discrete
- invariant under field transformations
- same in all frames

*Big bang singularity  
in Einstein frame is  
field singularity !*

$$g'_{\mu\nu} = \frac{\chi^2}{M^2} g_{\mu\nu} , \quad \varphi = \frac{2M}{\alpha} \ln \left( \frac{\chi}{\mu} \right)$$

choice of frame with constant particle masses is not well suited if physical masses go to zero !

# Crossover in quantum gravity and cosmology

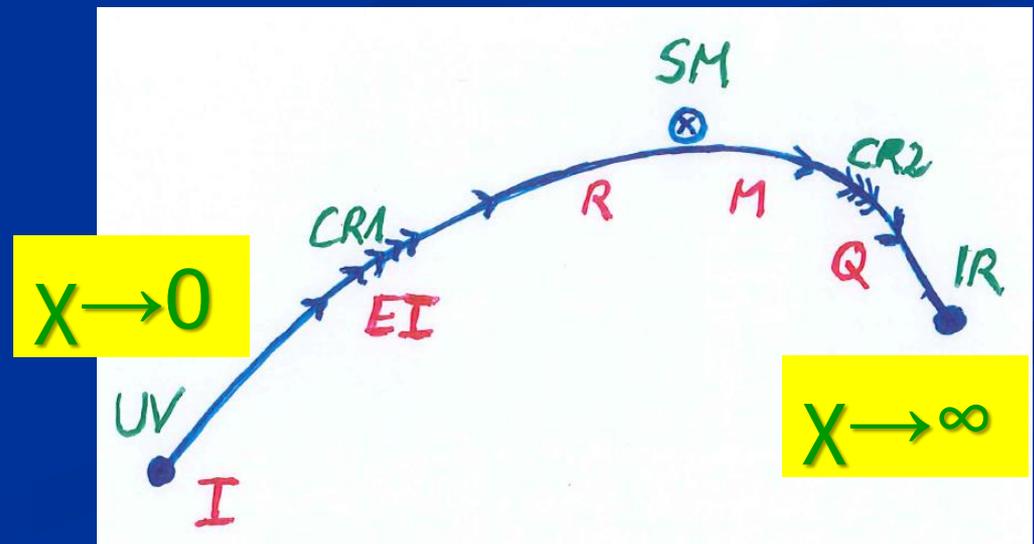


# Cosmological solution : crossover from UV to IR fixed point

- Dimensionless functions as  $B$  depend only on ratio  $\mu/\chi$ .
- IR:  $\mu \rightarrow 0$  ,  $\chi \rightarrow \infty$
- UV:  $\mu \rightarrow \infty$  ,  $\chi \rightarrow 0$

$$\Gamma = \int_x \sqrt{g} \left\{ -\frac{1}{2} \chi^2 R + \mu^2 \chi^2 + \frac{1}{2} (B(\chi/\mu) - 6) \partial^\mu \chi \partial_\mu \chi \right\}$$

Cosmology makes  
crossover between  
fixed points by  
variation of  $\chi$ .

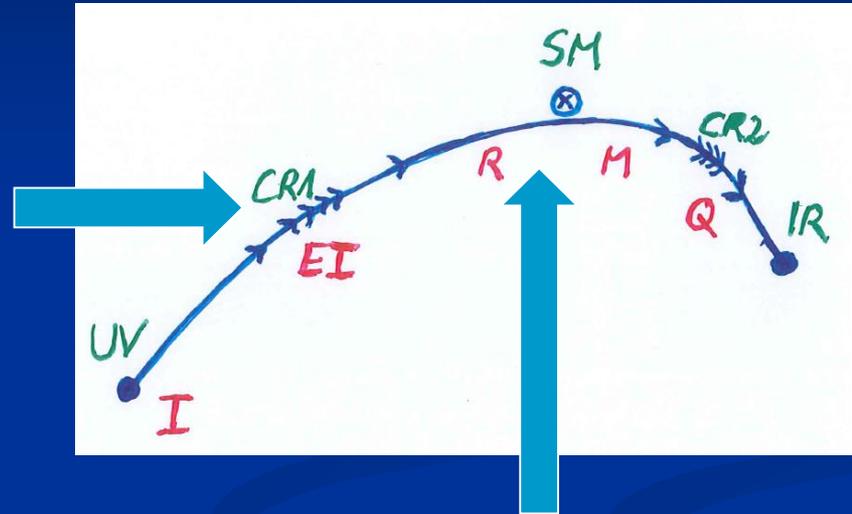


# Renormalization flow and cosmological evolution

- renormalization flow as function of  $\mu$   
is mapped by dimensionless functions to
- field dependence of effective action on scalar field  $\chi$   
translates by solution of field equation to
- dependence of cosmology on time  $t$  or  $\eta$

# Scaling solution

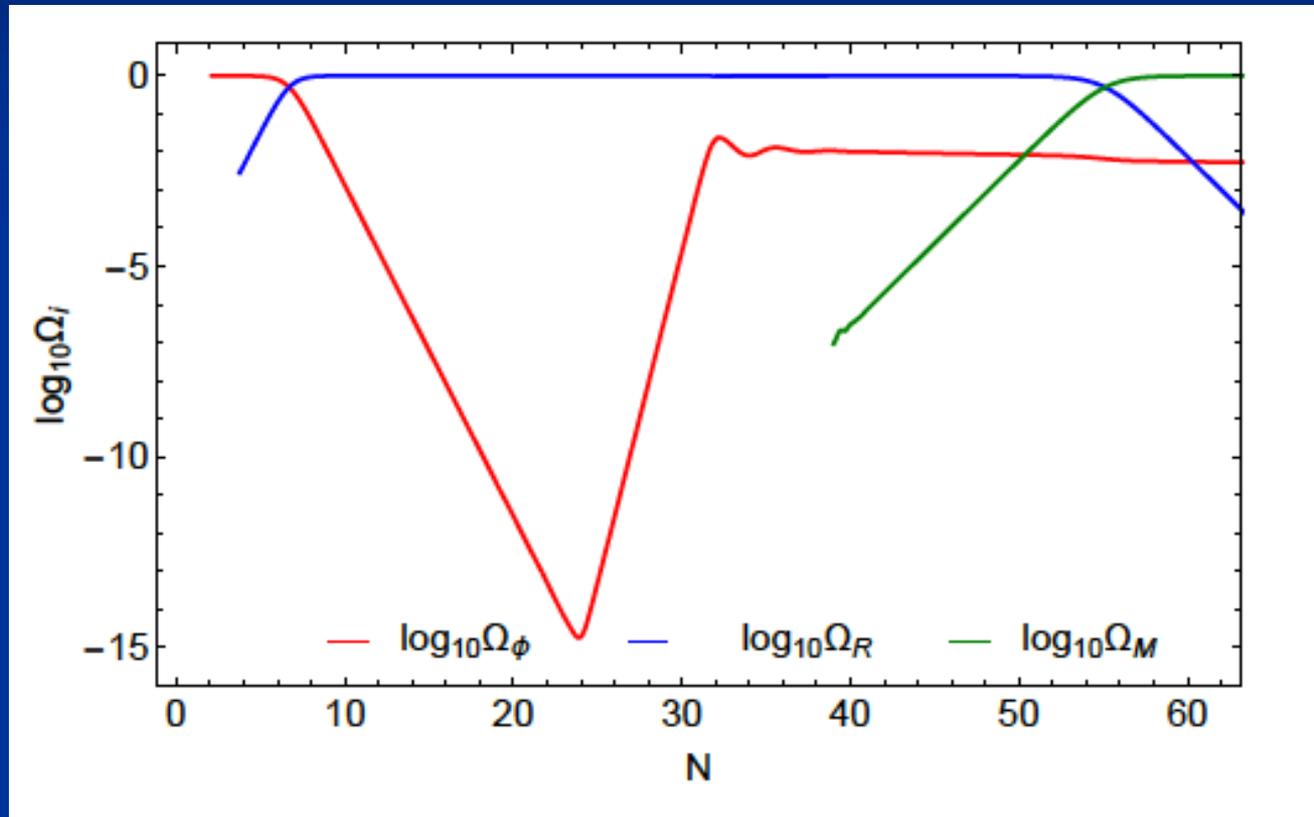
after end of inflation



scaling solution with  
few percent  
of Early Dark Energy

Dark Energy decreases  
similar to radiation and  
matter

# Evolution of dark energy fraction

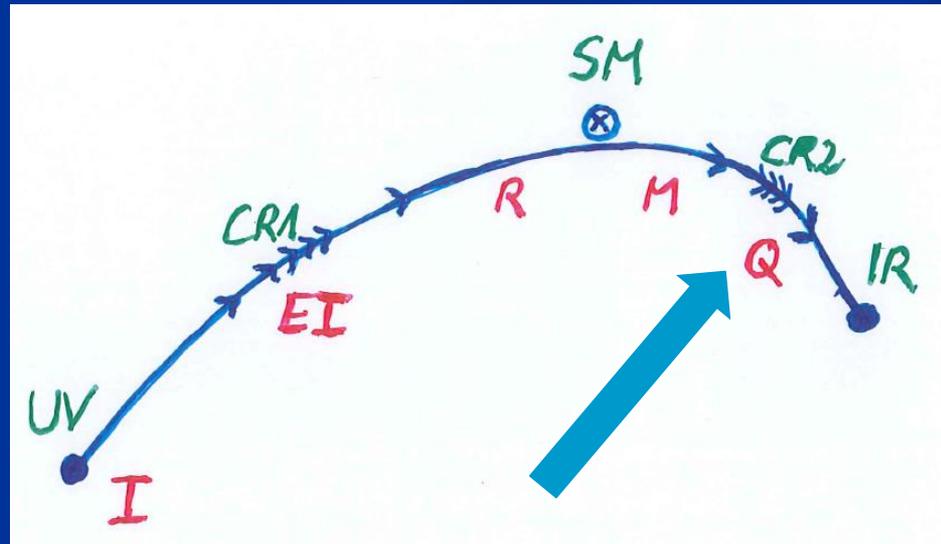


J. Rubio,...

# Growing neutrino masses and quintessence

# Second stage of crossover

- from SM to IR
- in sector Beyond Standard Model
- affects neutrino masses first ( seesaw or cascade mechanism )



# Varying particle masses at onset of second crossover

- **Except for neutrinos** all particle masses are proportional to  $\chi$ .
- Ratios of particle masses remain constant.
- Compatibility with observational bounds on time dependence of particle mass ratios.
- Neutrino masses show stronger increase with  $\chi$ , such that **ratio neutrino mass over electron mass grows**.

# Cosmic trigger

- Stop of evolution of scalar field when neutrinos become non-relativistic
- Transition from scaling solution to (almost) cosmological constant

# connection between dark energy and neutrino properties

$$[\rho_h(t_0)]^{\frac{1}{4}} = 1.27 \left( \frac{\gamma m_\nu(t_0)}{eV} \right)^{\frac{1}{4}} 10^{-3} eV$$

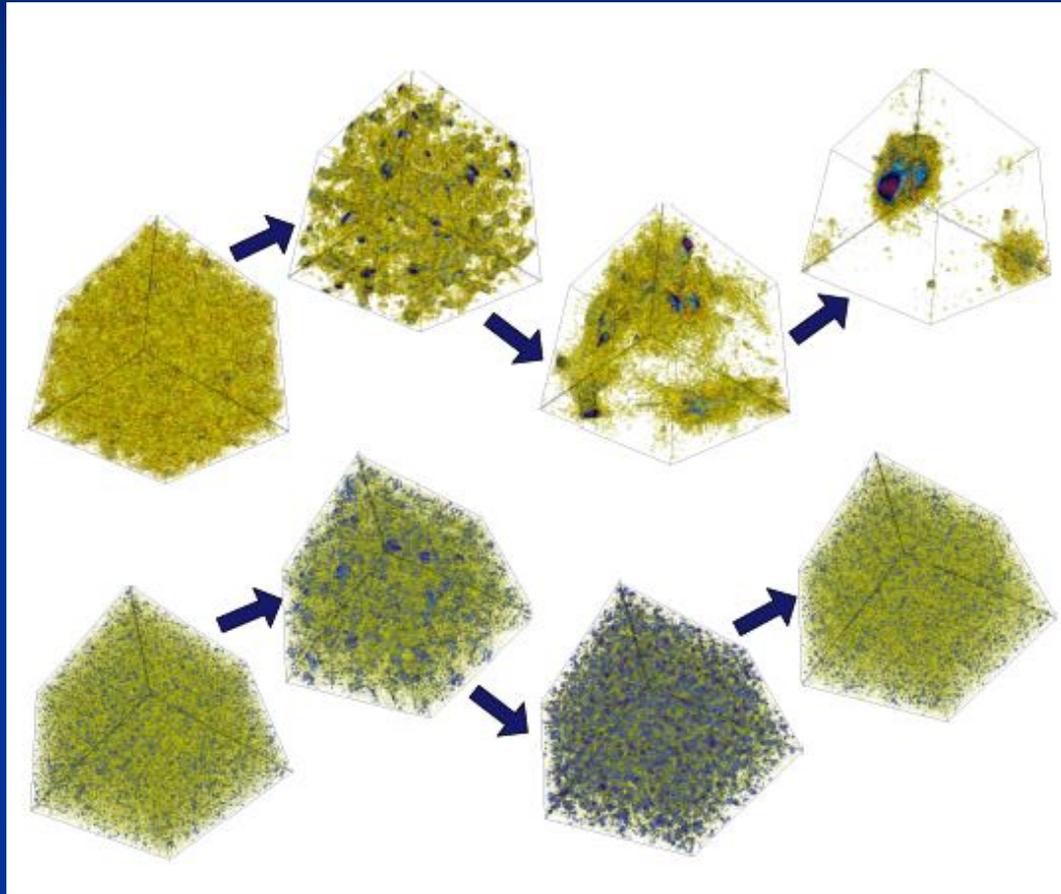
L.Amendola,  
M.Baldi, ...

present dark energy density given by neutrino mass

present equation  
of state given by  
neutrino mass !

$$w_0 \approx -1 + \frac{m_\nu(t_0)}{12eV}$$

# Neutrino lumps



large  $m_\nu$

small  $m_\nu$

Casas, Pettorino,...