Higher Moments of Net-Proton Multiplicity Distributions in Cu+Cu Collisions at $\sqrt{s_{NN}}$ = 22.4, 62.4 and 200 GeV from STAR

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Abstract

Fluctuation of conserved charges is important tool for studying the properties of QCD phase structures in high-energy nuclear collisions. Previously, the STAR experiment has reported the energy dependence of the cumulants of net-proton, net-charge and net-kaon distributions in Au+Au collisions at RHIC. Non-monotonic energy dependence has been observed in the 4th-order net-proton fluctuations in the top 5% central Au+Au collisions at RHIC. In this poster, we report the efficiency-uncorrected collision energy and centrality dependence of net-proton higher moments for Cu+Cu collisions at Vs_{NN}= 22.4, 62.4 and 200 GeV at RHIC.

Introduction

Efficiency Uncorrected Cumulants

- > Higher moments of the distribution of conserved quantities are predicted to be sensitive to the correlation length, and thus are suitable for the search for the QCD critical point^[1,2,3].
- > Main observables:

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Volume independent cumulant ratios:

 $\kappa \sigma^2 = \frac{C_4}{C_2} = \frac{\chi_4}{\chi_2}, \ S\sigma = \frac{C_3}{C_2} = \frac{\chi_3}{\chi_2}, \ \frac{\sigma^2}{M} = \frac{C_2}{C_1} = \frac{\chi_2}{\chi_1}$

 \succ Cumulant ratios of Cu+Cu collisions from RHIC similar to those of the Au+Au collisions present a new opportunity to measure the system size dependence^[4].

Particle Identification







The cumulants of proton, anti-proton and net-proton multiplicity distributions linearly increase with the number of participant.



 \succ The ratios of proton, anti-proton are close to unity.

Summary

- The efficiency-uncorrected cumulants and their ratios of netproton have been measured in Cu+Cu collisions at V_{NN} =22.4, 62.4 and 200 GeV.
- The values of $C_2/C_1(C_3/C_2)$ of net-proton increase (decrese)

with increasing collision energy, while the C_4/C_2 values are flat as a function of energy.

Efficiency corrections and systematic uncertainty estimation are in progress.



1.S.Gupta, X.Luo, B.Mohanty, H.G.Ritter, N.Xu, Science 332, 1525 (2011). 2.L.Adamczyk et al., (For the STAR Collaboration), Phys.Rev.Lett.113(2014) 92301. 3.L.Adamczyk et al., (For the STAR Collaboration), Phys.Rev.Lett.112(2014)32302. 4.B.I.Abelev et al., (For the STAR Collaboration), Phys.Rev.C 81, 044902 (2010).



The STAR Collaboration drupal.star.bnl.gov/STAR/presentations

