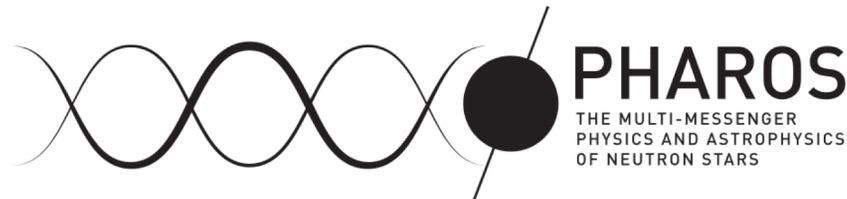


Signatures of Quark-Hadron Phase Transitions in General-Relativistic Neutron-Star Mergers

Veronica Dexheimer

in collaboration with Elias Most, Jens Papenfort,
Matthias Hanauske, Luciano Rezzolla and Horst Stöcker

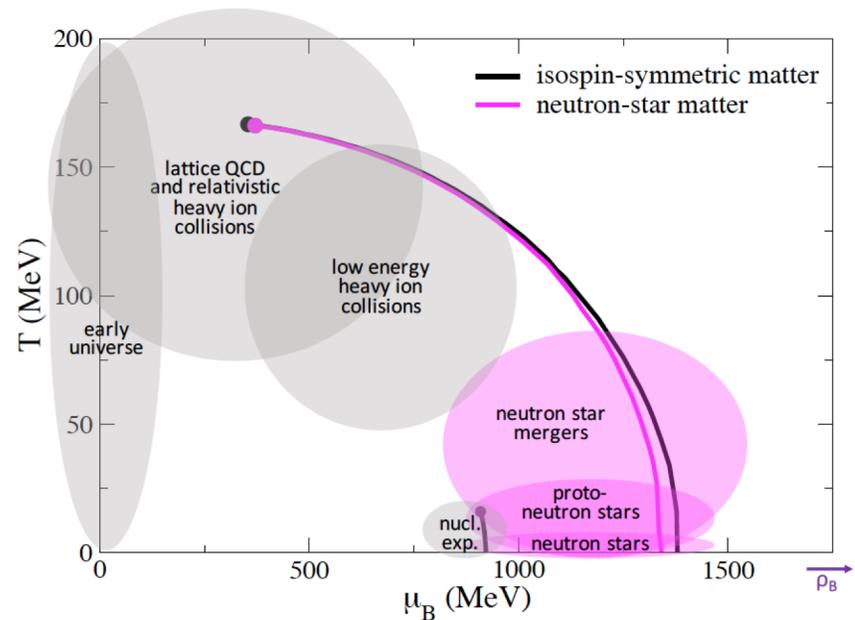
Phys. Rev. Lett. (2019)
J. Phys. G (2019)



EoS Ingredients for NS Mergers

(in addition to nuclear and astrophysical constraints)

- high density description
- quantum relativistic description
- finite/high temperature description
- provides particle population
- include hyperons and quarks
- include chiral symmetry restoration
- in agreement with lattice QCD
- in agreement with heavy-ion collision results
- in agreement with perturbative QCD results



CMF (Chiral Mean Field) Model

- non-linear realization of the linear sigma model
- includes baryons (+ leptons) and quarks
- fitted to reproduce nuclear, astrophysical, lattice QCD, and heavy ion, constraints
- baryon and quark effective masses

$$M_B^* = g_{B\sigma}\sigma + g_{B\delta}\tau_3\delta + g_{B\zeta}\zeta + M_{0_B} + g_{B\Phi}\Phi^2$$
$$M_q^* = g_{q\sigma}\sigma + g_{q\delta}\tau_3\delta + g_{q\zeta}\zeta + M_{0_q} + g_{q\Phi}(1 - \Phi)$$

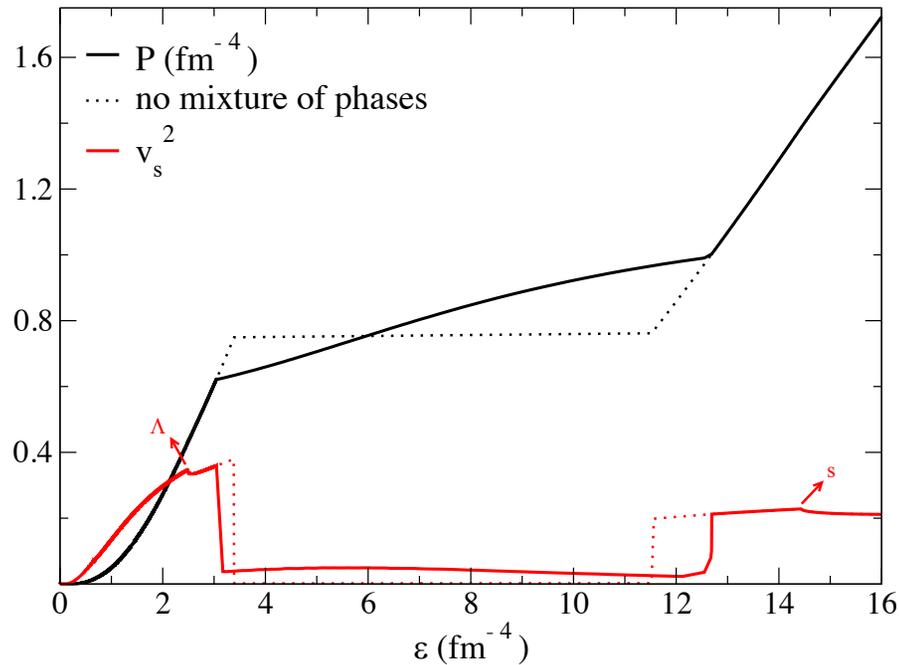
- 1st order phase transitions or crossovers

- potential for Φ
deconfinement
order parameter

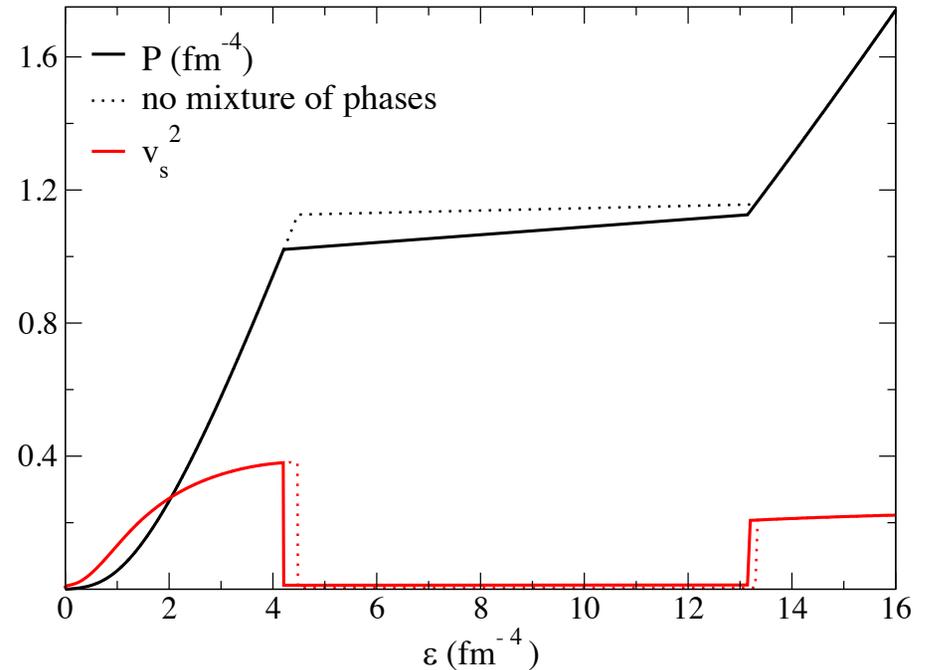
$$U = (a_0 T^4 + a_1 \mu_B^4 + a_2 T^2 \mu_B^2) \Phi^2 + a_3 T^4 \ln(1 - 6\Phi^2 + 8\Phi^3 - 3\Phi^4)$$

Speed of sound for neutron-star matter

T=0



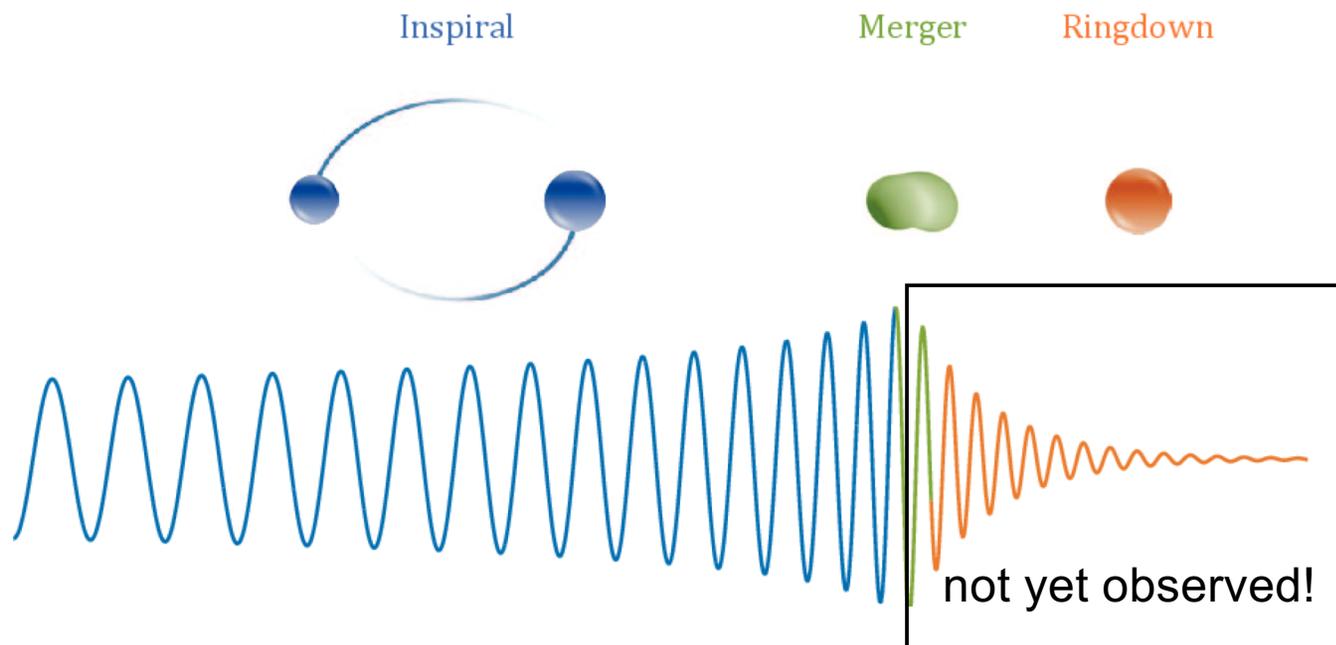
$S_B=2$ ($Y_I=2$)



- $v_s^2 \sim 1/3$ at very large energies
- tested against PQCD results at zero/finite temperature

Neutron Star Merger 170817

- gravitational waves observed by LIGO/VIRGO in 17 August 2017 from galaxy NGC 4993 140 million light-years away
- observed electromagnetically by 70 observatories on 7 continents and in space

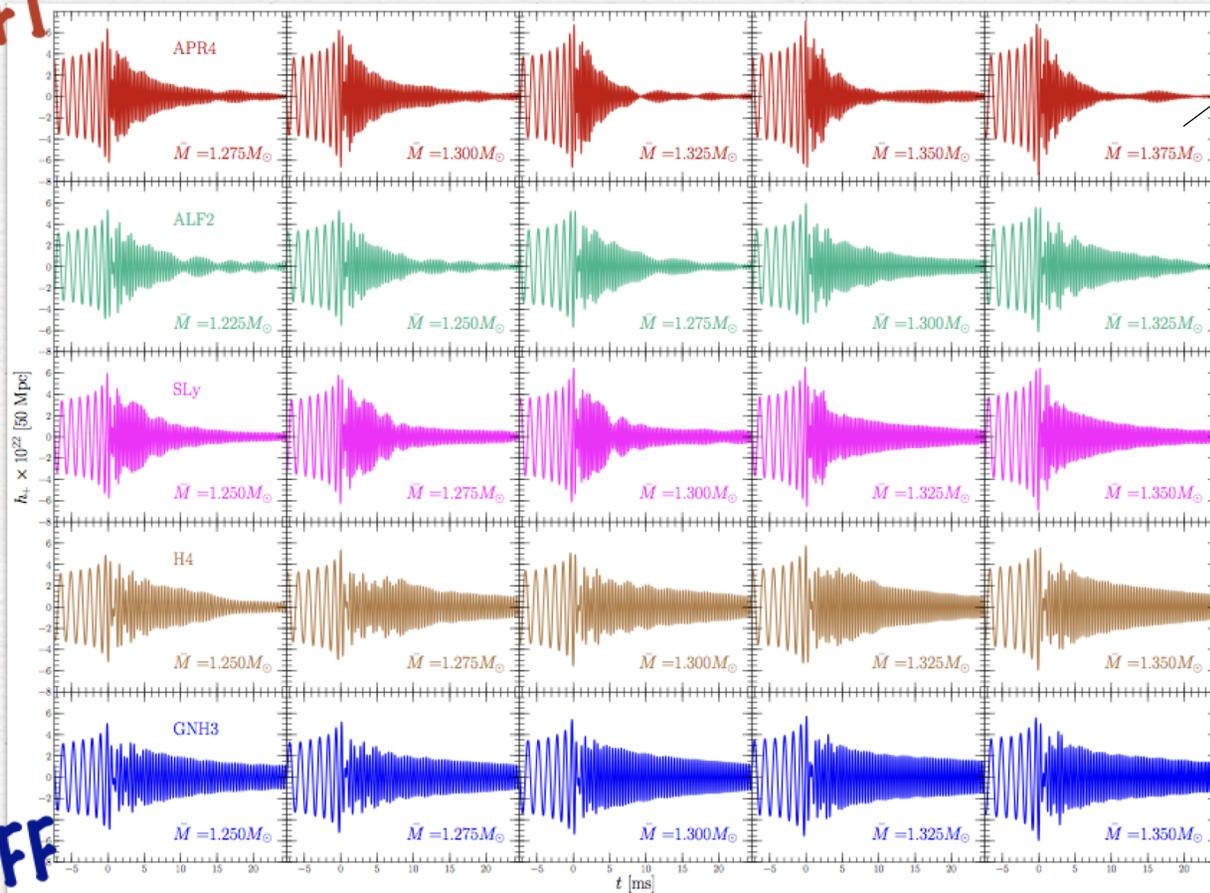


Hadronic NS Merger Simulations

What we can do nowadays

Takami, Rezzolla, Baiotti (2014, 2015), Rezzolla+ (2016)

SOFT



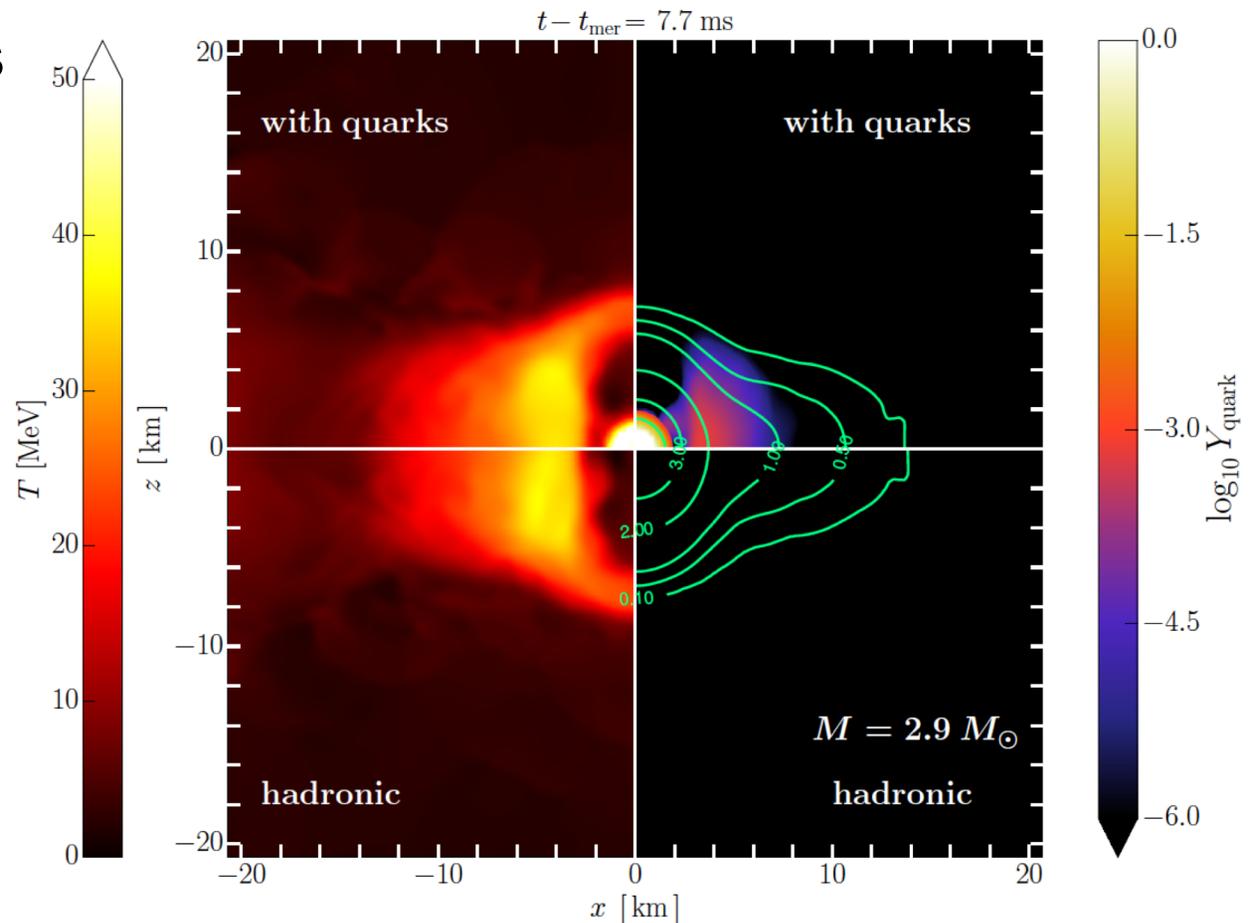
average stellar mass

STIFF

NS Merger Simulation with Deconfinement

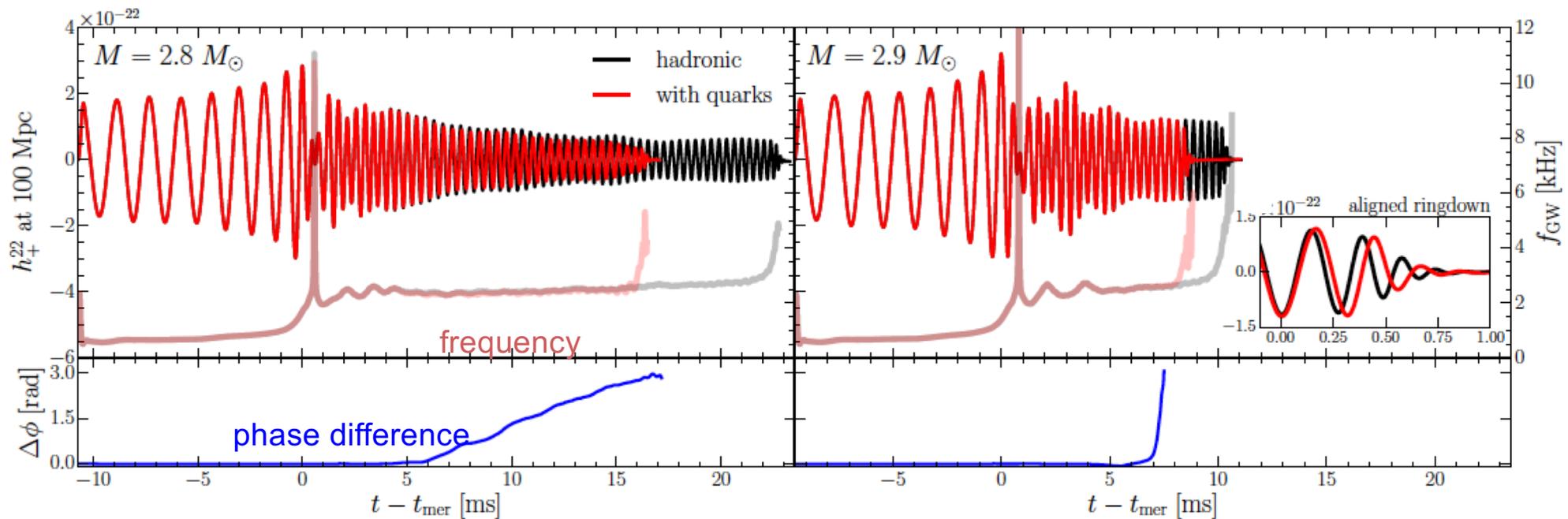
- 3D (T, n_B, Y_c) CMF EoS with/without quarks
- solve coupled Einstein-hydrodynamics system using Frankfurt/IllinoisGRMHD code (FIL)

- interesting results for binaries with final masses of 2.8 and $2.9 M_{\text{Sun}}$ after deconfinement but before collapse to black hole



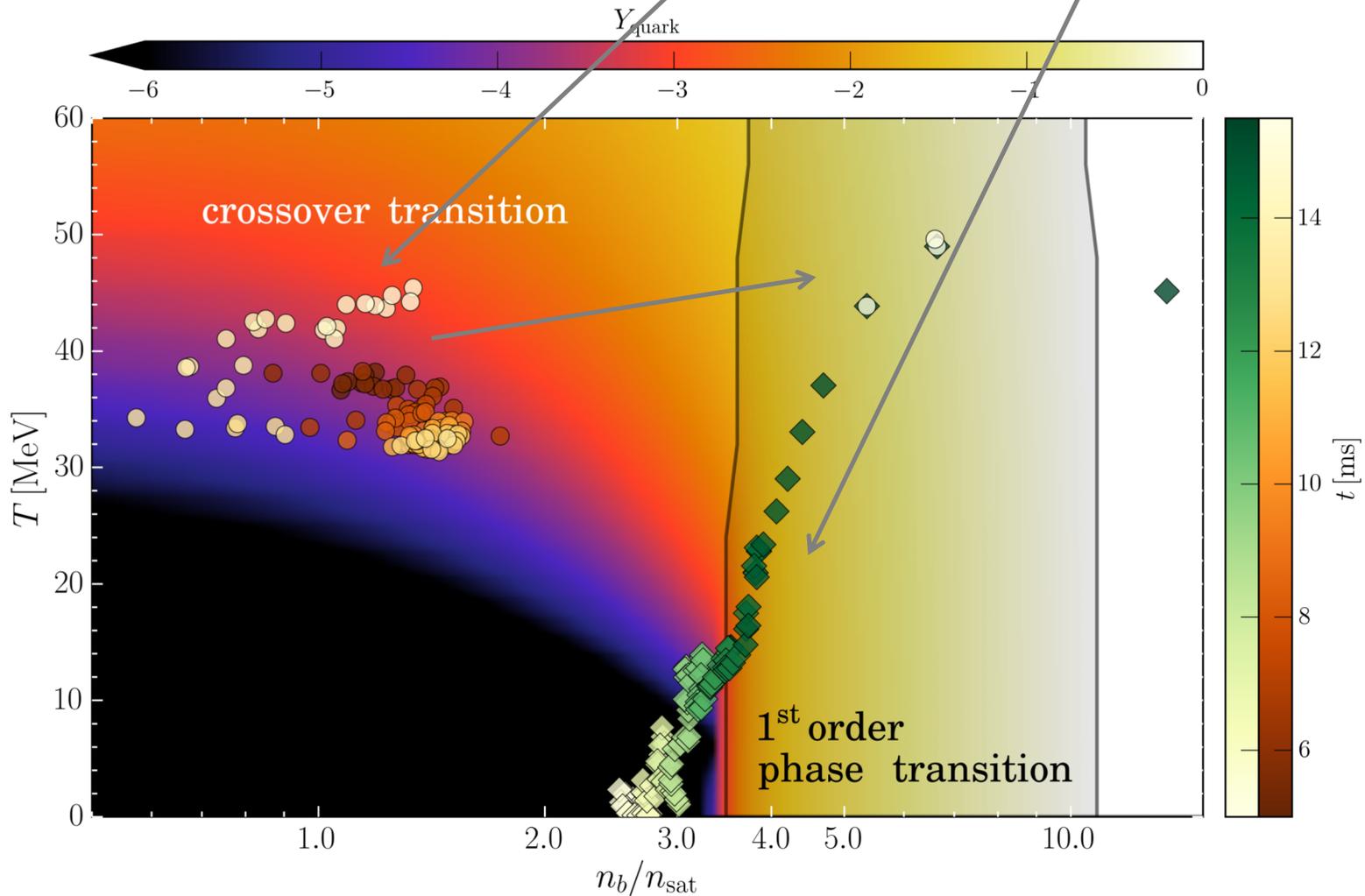
NS Merger Simulation with Deconfinement

- effects from quarks only after the merger
- change in frequency, strain, and time for collapse to blackhole
- collapse is faster in more massive case



Merger in the QCD phase diagram:

- Tracking maximum temperature ● and density ◆ in merger



Tidal Deformability

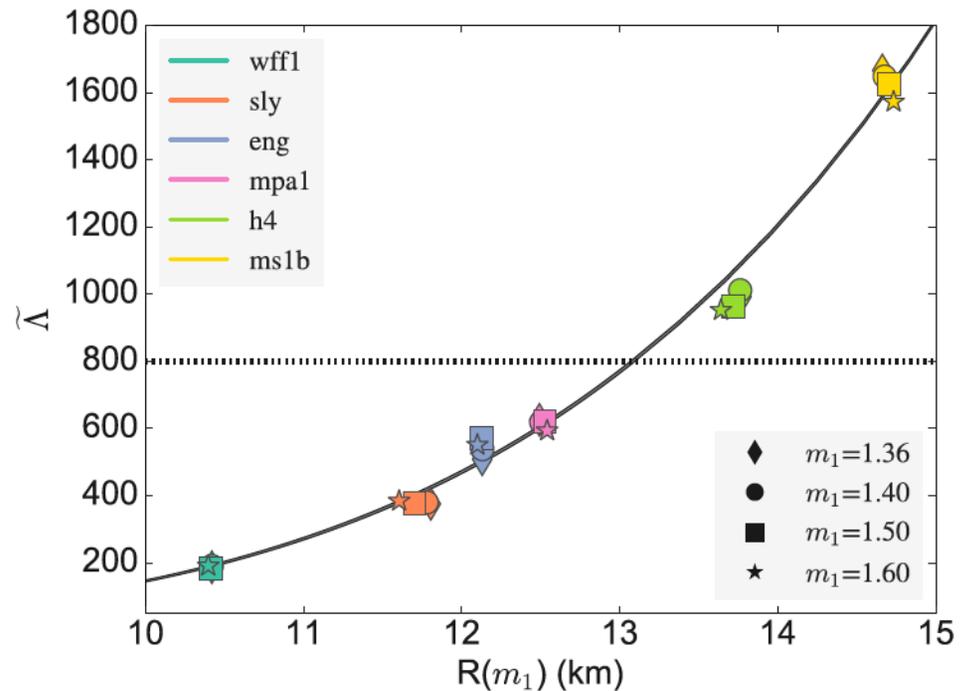
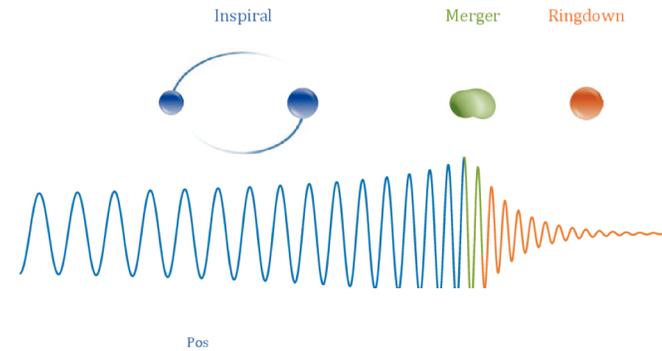
- normalized stellar quadrupole deformation by companion

- calculated from finite-size effects in end of inspiral:

76 \rightarrow 1045 with 90% confidence (De et. al 2018)

- related to NS radius of $M=1.4 M_{\text{sun}}$ (Raithel et. al 2018)

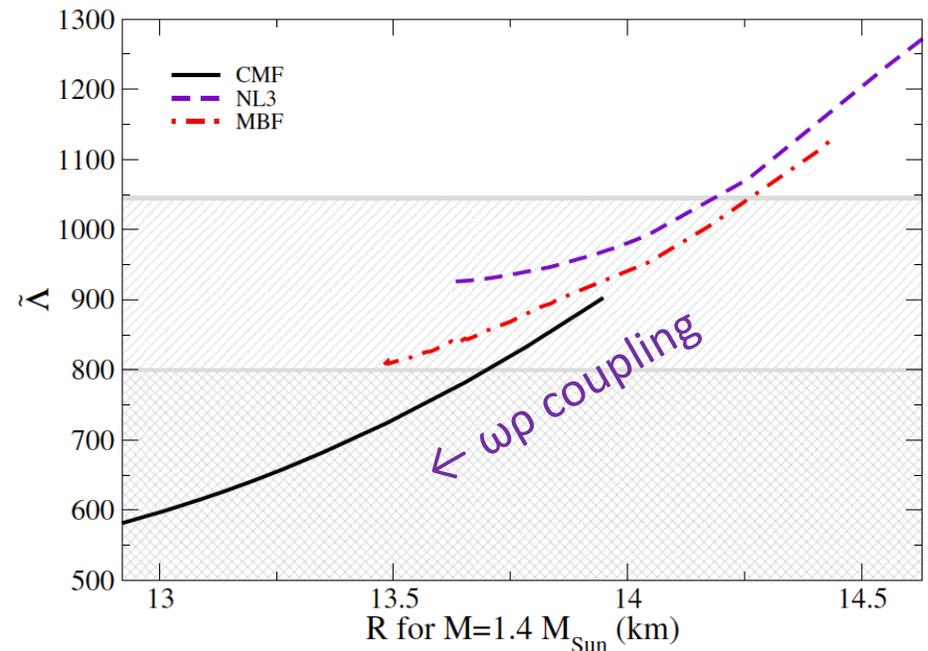
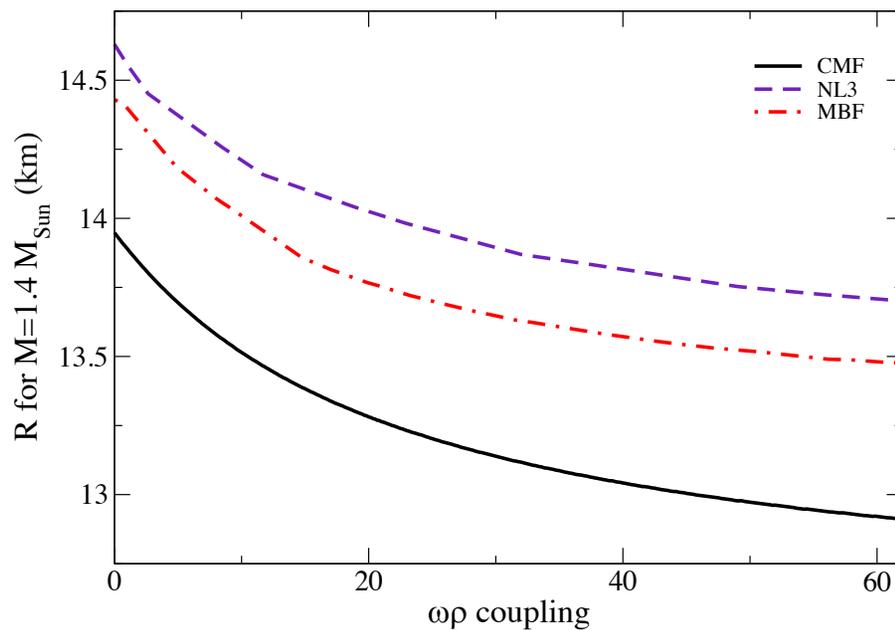
- universal relation?



Exploring hadronic isovector coupling

- using 3 relativistic EoS's that fulfil standard nuclear and astrophysical constraints: NL3, MBF, and CMF

- new repulsive vector-isovector channel $L_{\omega\rho} = g_{\omega\rho} g_{\omega}^2 g_{\rho}^2 \omega_{\mu} \omega^{\mu} \rho_{\mu} \rho^{\mu}$ suggested by Horowitz and Piekarewicz (2002) $\omega\rho$ coupling



- non-trivial relation between $\tilde{\Lambda}$ and $R_{1.4M_{\text{Sun}}}$

Outlook

- GW170817 provided us a new way to study dense and hot matter
- more quantum relativistic models with temperature/exotic degrees of freedom needed to study
 - relation between tidal deformability and nuclear physics
 - realistic neutron-star merger simulations
- more merger data coming soon ...so, maybe, a clear first signature for quark deconfinement phase transition will come from astrophysics?