Neutrinos from Gamma-Ray Bursts and Quasars

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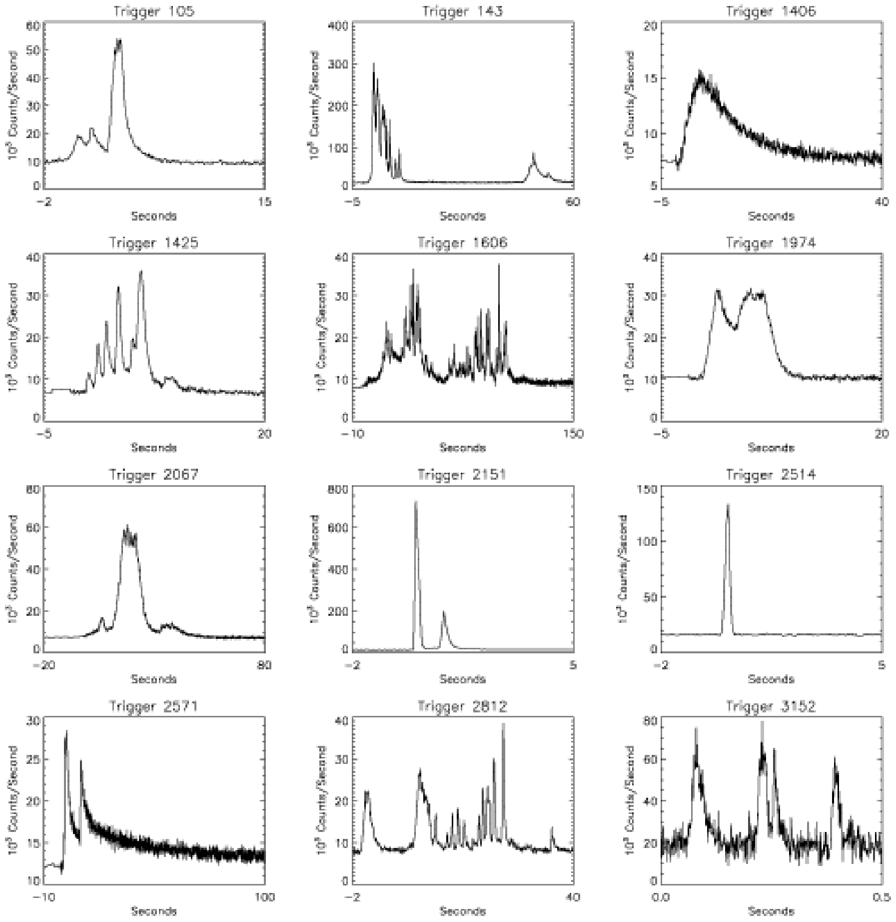
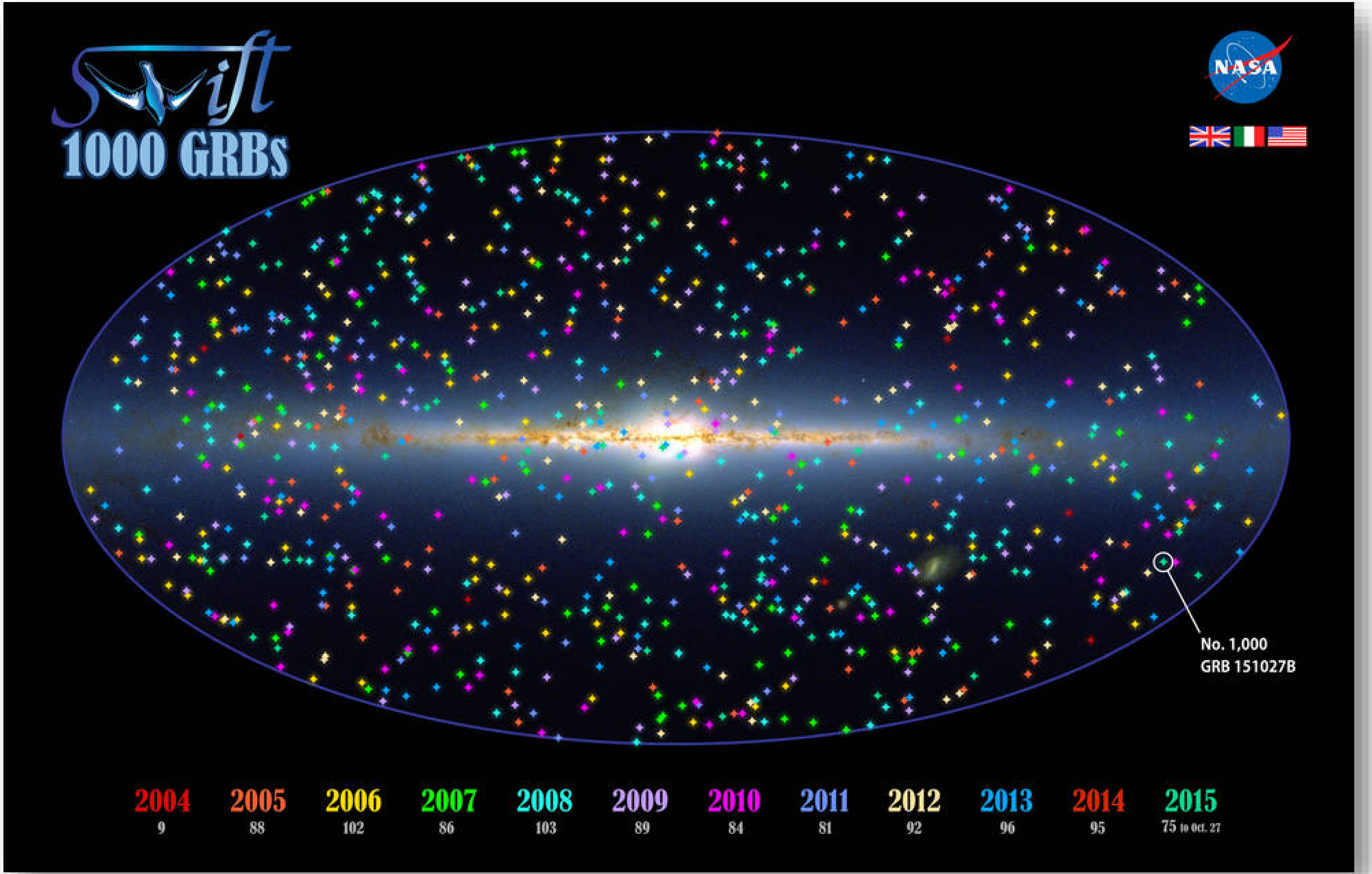
Gamma-Ray Bursts (GRBs)

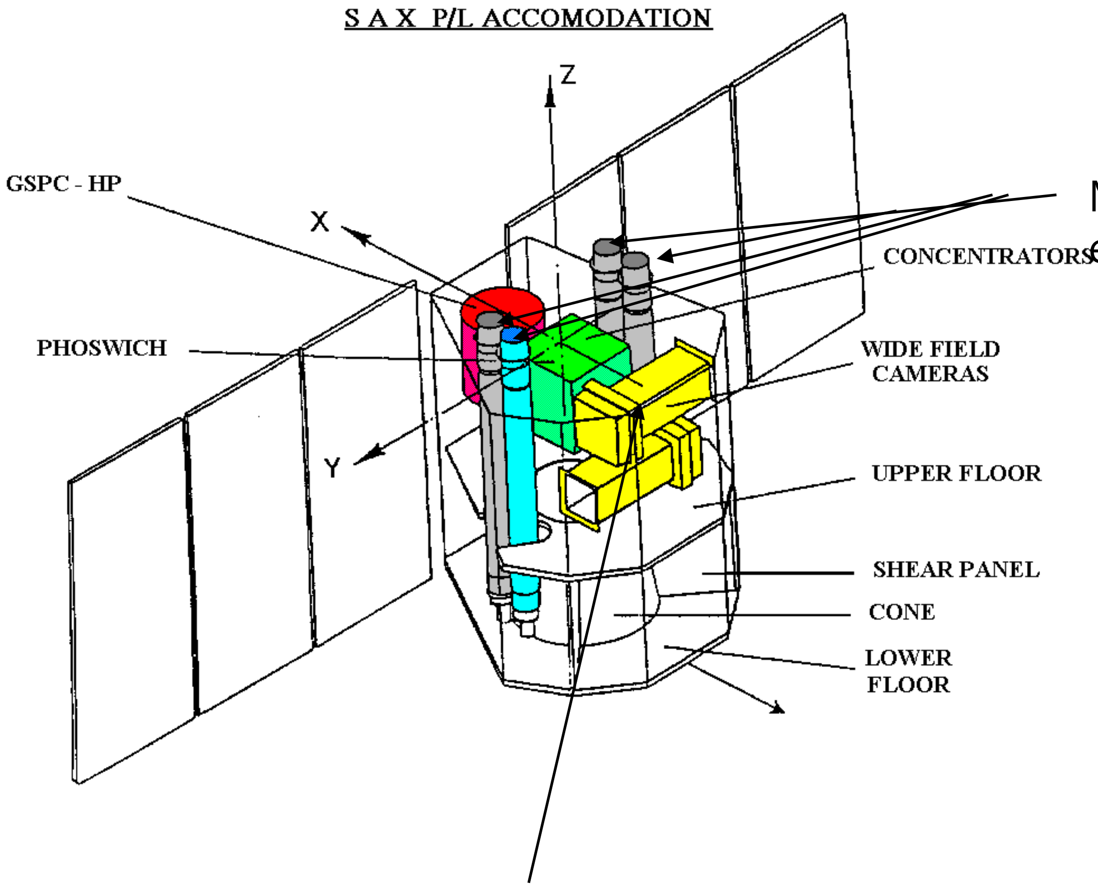
- Brief, bright flash of gamma-rays
- Vela 5 (1969~1979) detected 73 GRBs
- Timescales of 10 msec ~ hours
- Wide variation both in time-structure and duration
- As bright as the visible stars, some as bright as Venus
- Never repeat from the same source



Identification of GRBs

- Isotropic distribution in the sky
  - Local: Solar System, Galactic Halo?
  - Cosmological
  - 1973 ~ 1993
- Compton Gamma-Ray Observatory (1991~2000)
- BeppoSax (1996~2002)
  - Italian-Dutch X-ray mission
  - 0.1~200 keV, good energy resolution
  - Wide-field cameras
    - Capable of monitoring X-ray transient phenomena
  - 1 msec resolution from 60~600 keV


Swift
1000 GRBs
NASA
No. 1,000
GRB 151027B
2004
9
2005
88
2006
102
2007
86
2008
103
2009
89
2010
84
2011
81
2012
92
2013
96
2014
95
2015
75 to Oct. 27



MEC S and LECS (medium and low energy x-ray sensors, 1 arc min positions)

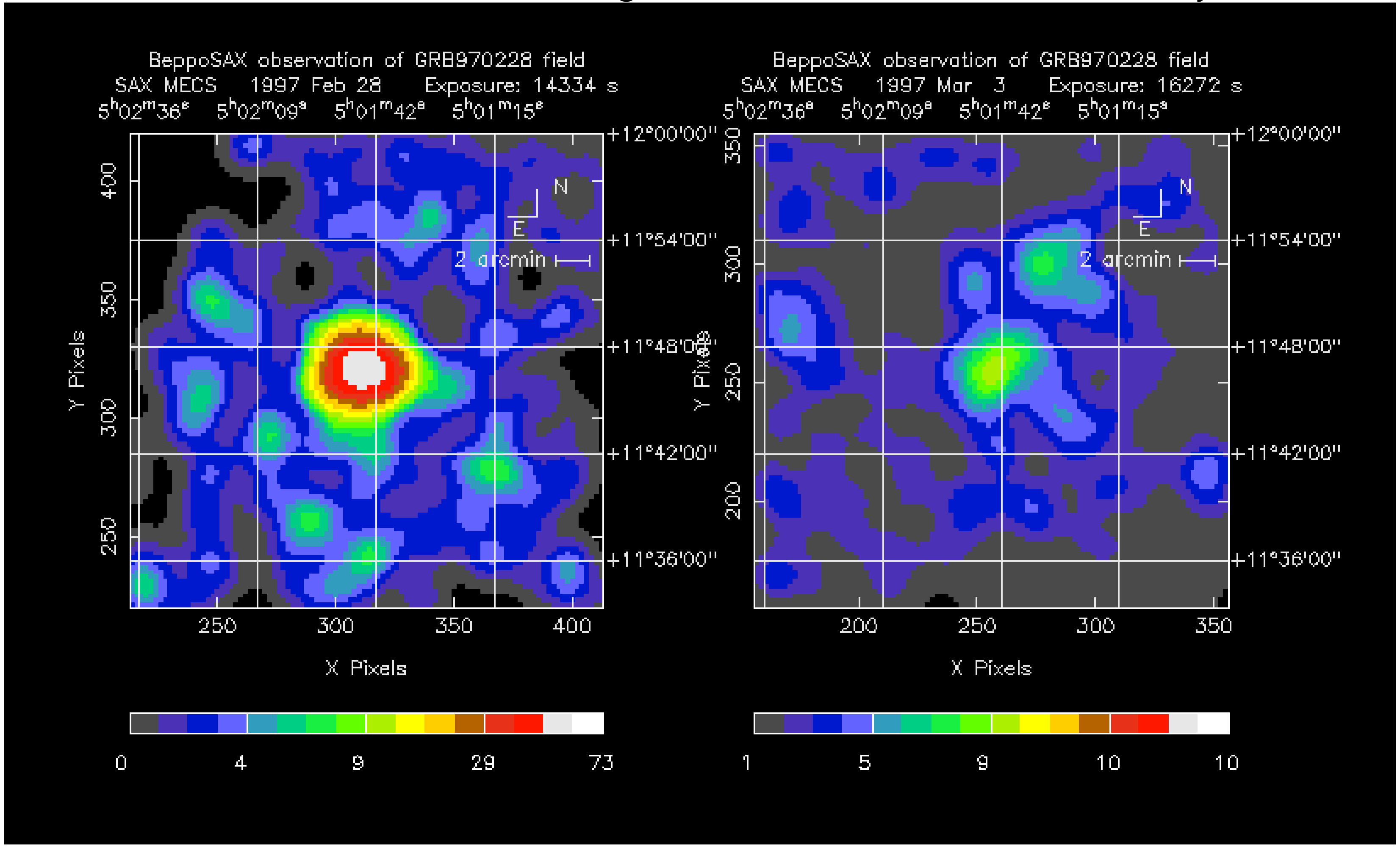
(2-30 keV; 20x20 degree FOV angular resolution 5 arc min)

The scintillator anti-coincidence shields of the Phoswich detector are able to detect gamma-rays 60-600 keV and get crude angular information

Woosley

BeppoSax GRB 970228 (discovered with WFC)

Feb 28, 1997 (8 hr after GRB using MECS)   March 3, 1997 (fainter by 20)

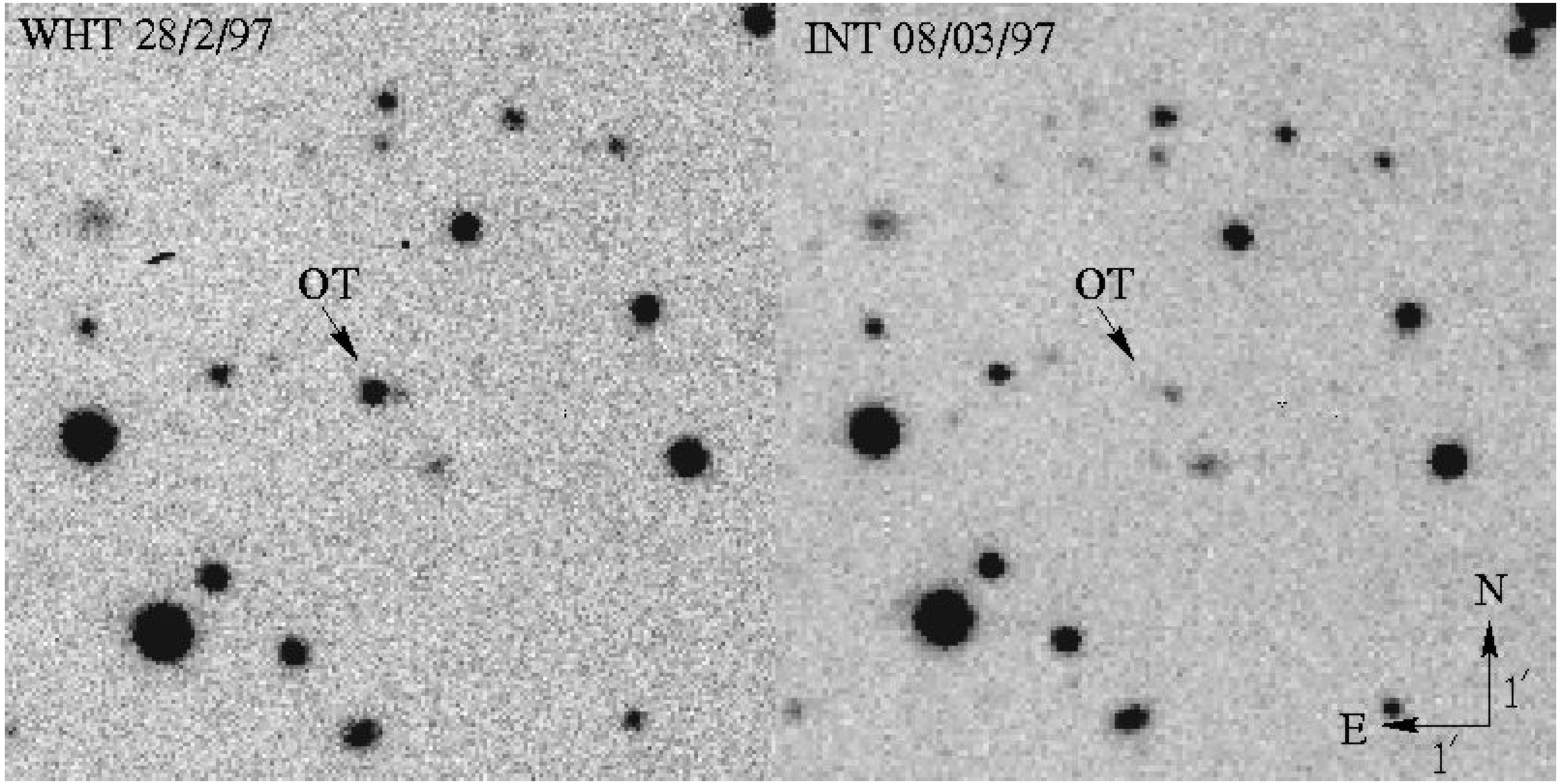


Each square is about 6 arc min

Woosley

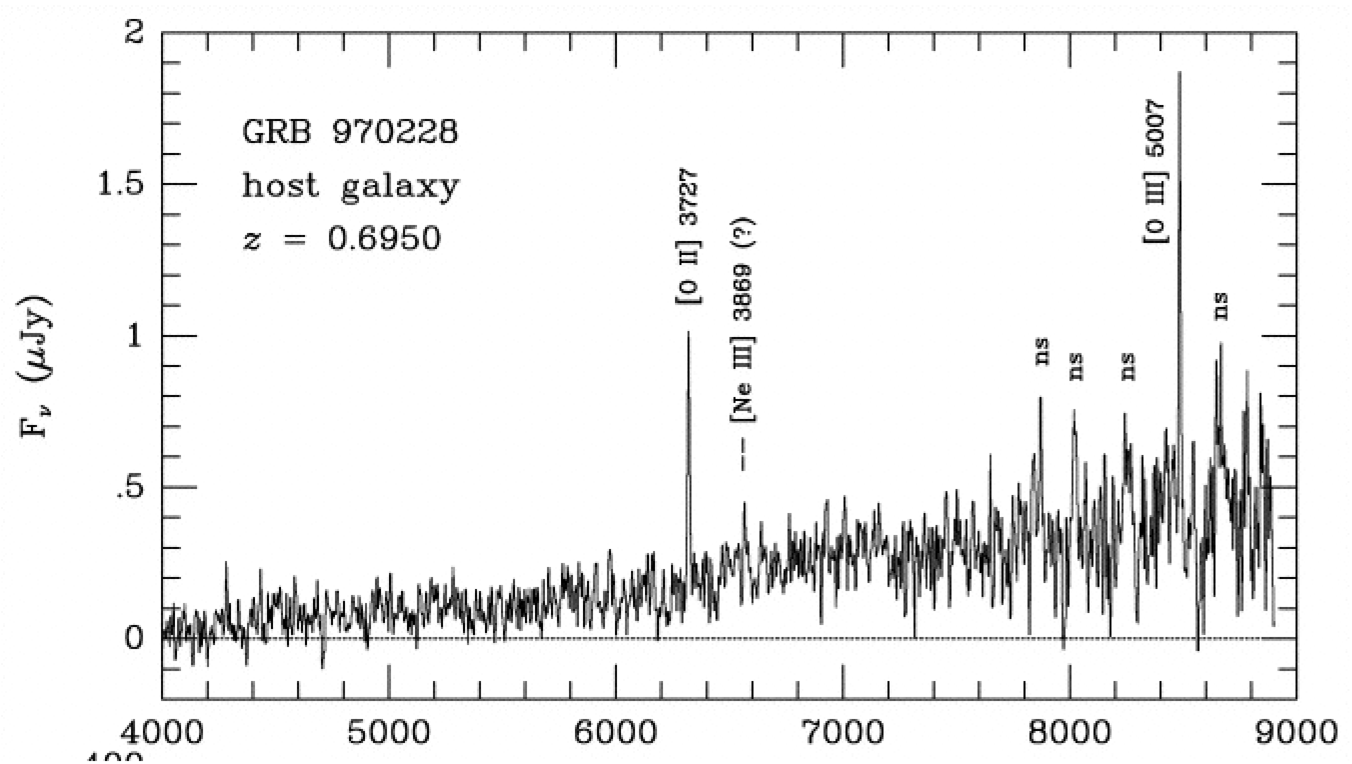
GRB 970228

William Hershel Telescope　　Isaac Newton Telescope



Groot, Galama, von Paradijs, et al IAUC 6584, March 12, 1997

Woosley



Spectrum of the host galaxy of GRB 970228 obtained at the Keck 2 Telescope. Prominent emission lines of oxygen and neon are indicated and show that the galaxy is located at a redshift of z = 0.695. (Bloom, Djorgovski, and Kulkarni (2001), ApJ, 554, 678. See also GCN 289, May 3, 1999.

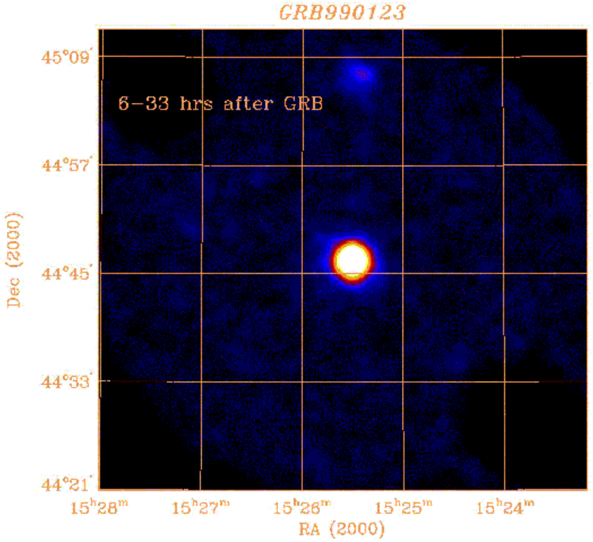
Energetics of GRB 970228

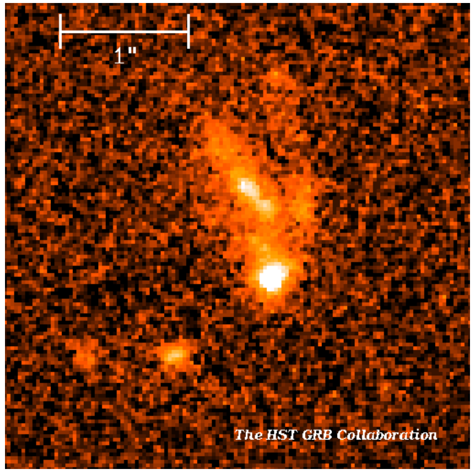
- Distance
  - z = 0.6950
  - Cosmological distance ~ 4.274 Gpc
- Brightness
  - 1.6 x 10^{52} erg in gamma-rays alone
  - ~ 2000 x L_{SN}
    - isotropic emission assumed

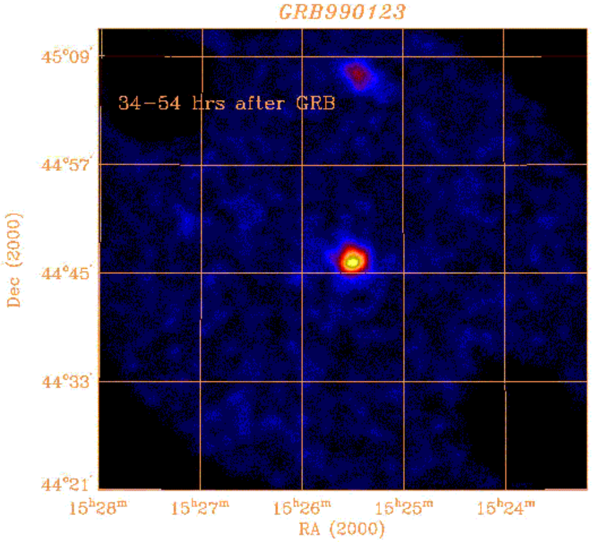
GRB 990123

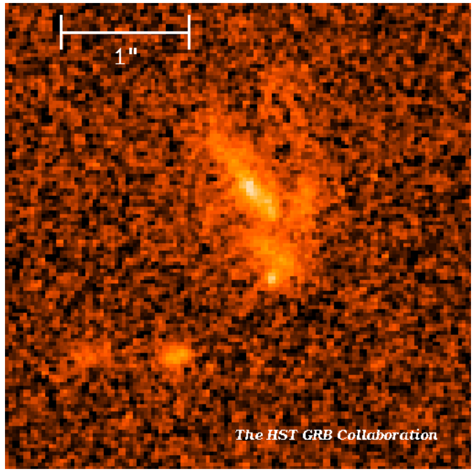
- $z = 1.61$
- $E \sim 10^{54}$ erg

- Typical GRBs
  - $z \sim 1$
  - $E \sim 10^{53}$ erg



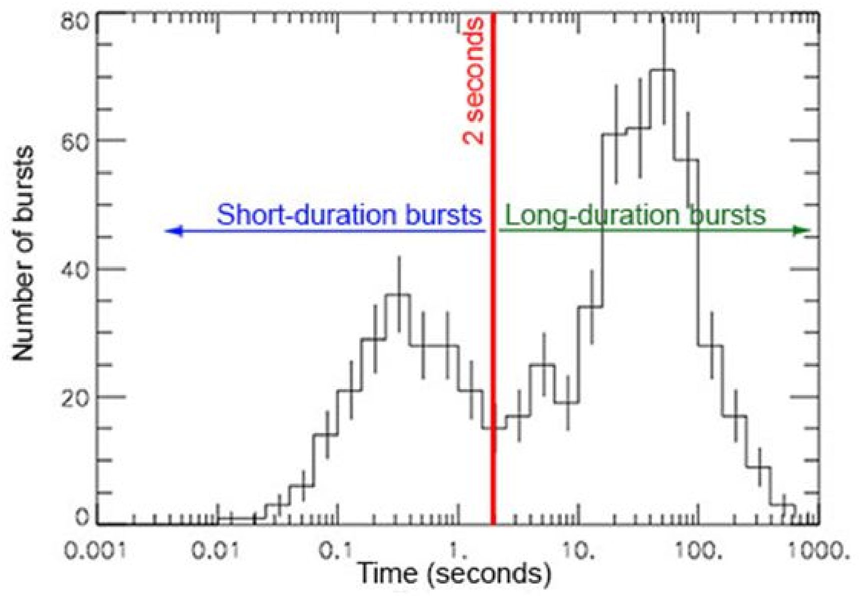






Short and Long GRBs

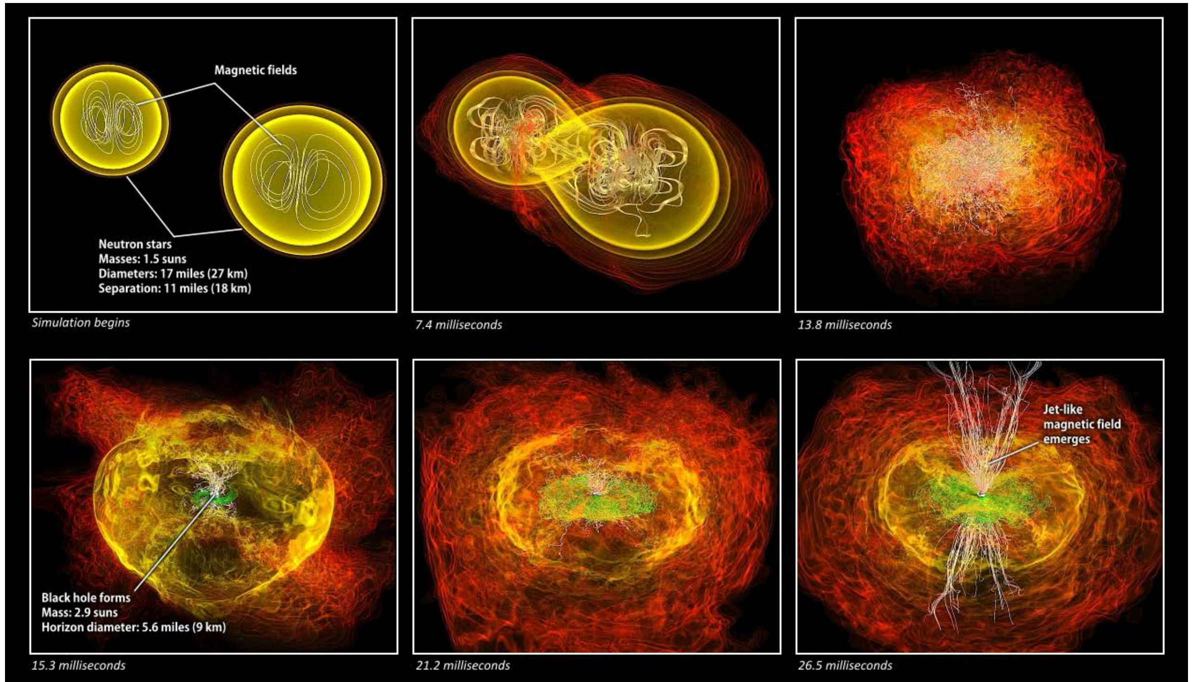
- Short vs Long GRBs
  - Short GRBs (SGRBs): t < 2 sec
  - Long GRBs (LGRBs): t > 2 sec
- LGRBs
  - ~ 70% of GRBs
  - Associated with star-forming regions
  - Many associated with a core-collapse supernova
- SGRBs
  - ~ 30% of GRBs
  - Not associated with star formation
  - Associated with elliptical galaxies, the central regions of large galaxy clusters



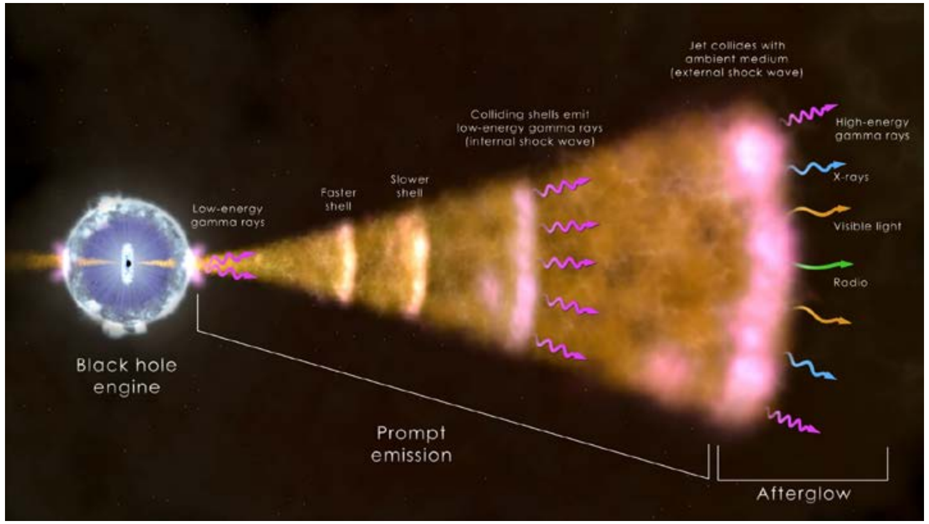
Graph of the time versus number of bursts for the gamma-ray bursts observed by the BATSE instrument on the Compton Gamma-ray Telescope.

Origin of GRBs

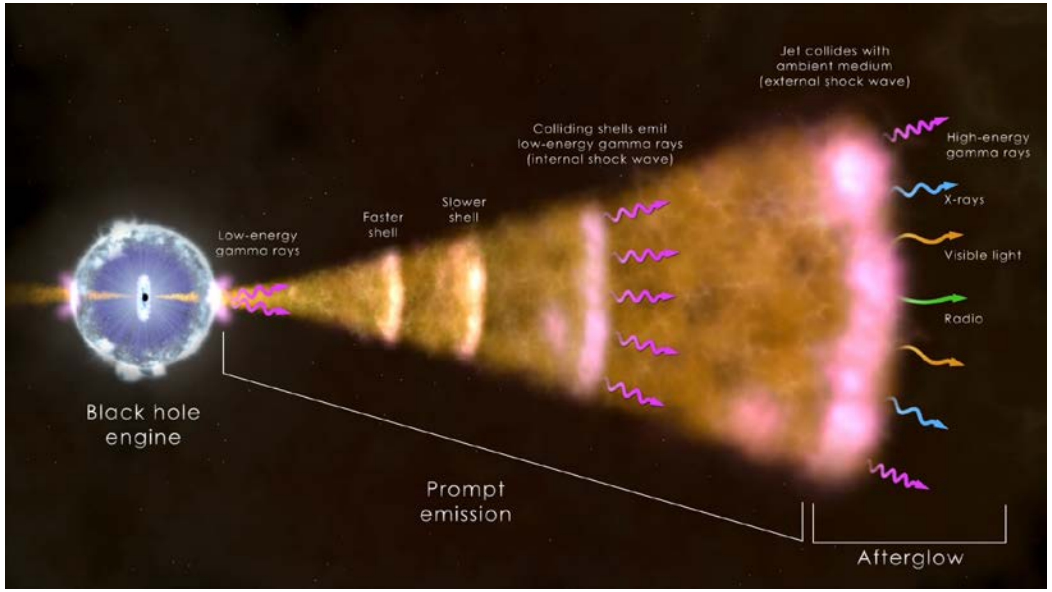
- Short GRBs
  - Compact GR emission region
  - NS-NS or BH-NS (BNS) merger model
    - GW170817
      - GW signal from NS-NS merger
      - GRB 170817A – short GRB
      - Afterglow - kilonova

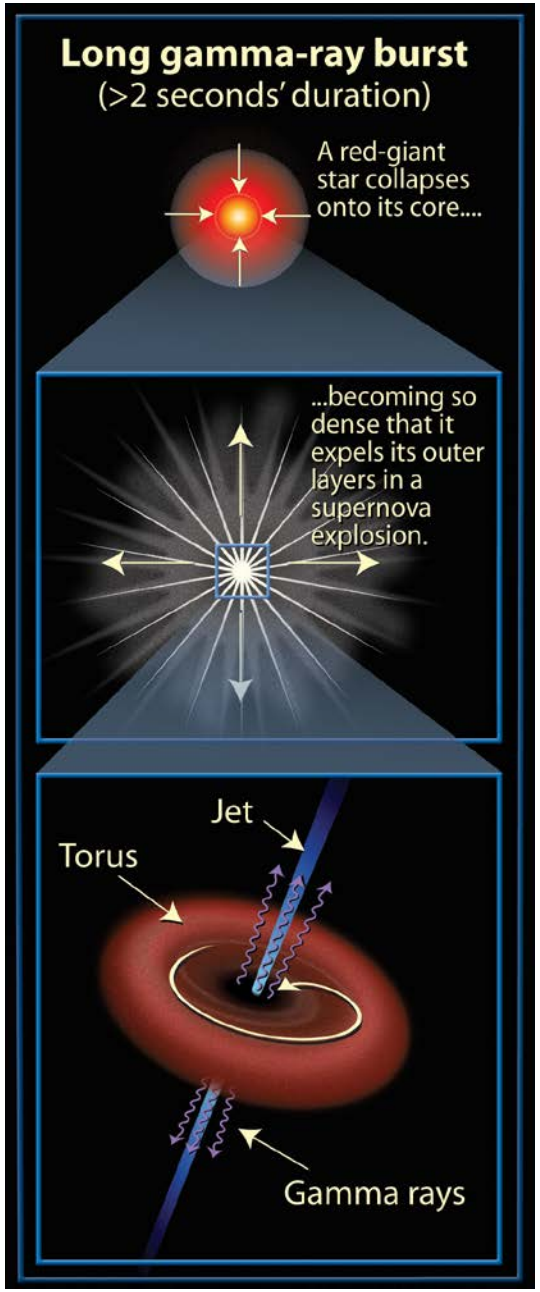


Credit: NASA/AEI/ZIB/M. Koppitz and L. Rezzolla



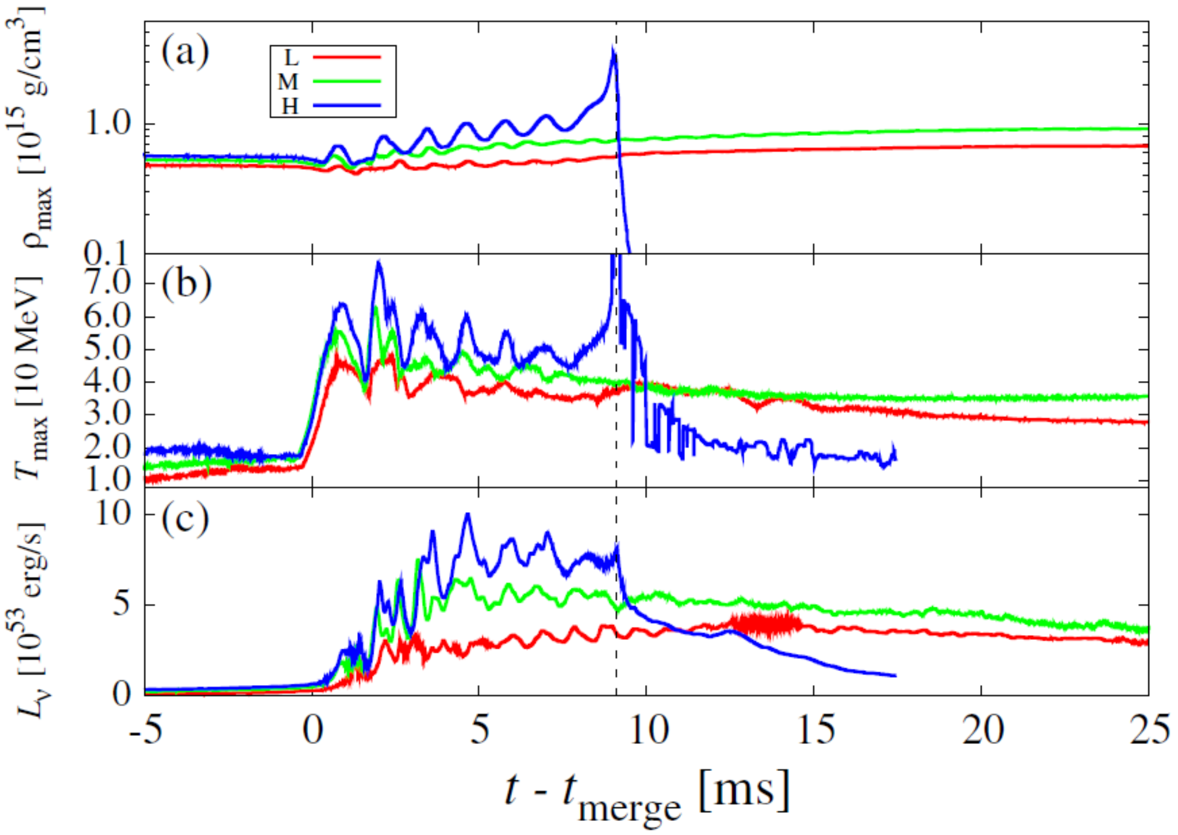
- Long GRBs
  - Collapsar model
    - Core of an extremely massive, low-metallicity, rapidly rotation star collapses into a black hole





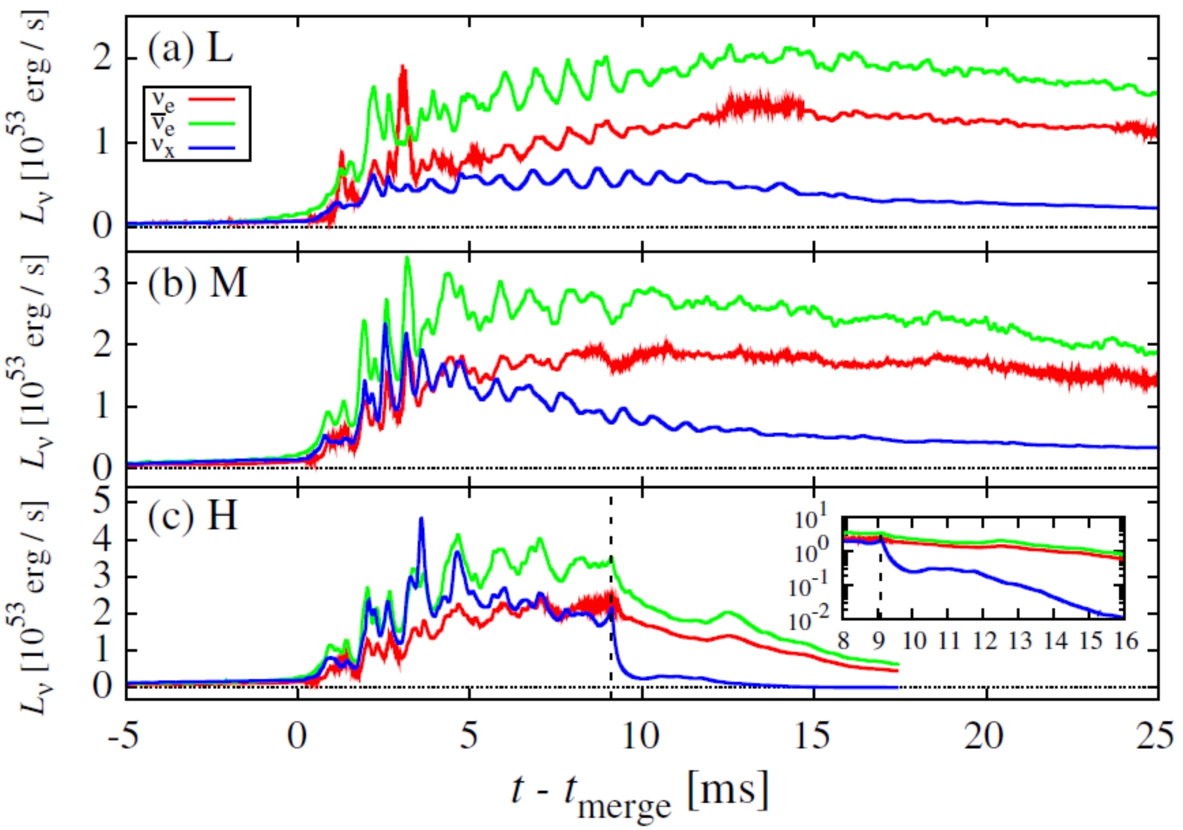
NS-NS Merger

- Gravitational binding energy should be released as in supernovae
    - ~ 3 x 10^{53} erg
    - Released as thermal and kinetic energy
    - Some of these energy in the form of photons and neutrinos
- Relativistic radiation hydrodynamic simulations (Sekiguchi+ 2011)
    - ~ 3-8 x 10^{53} erg/s
    - 20-75 MeV neutrinos

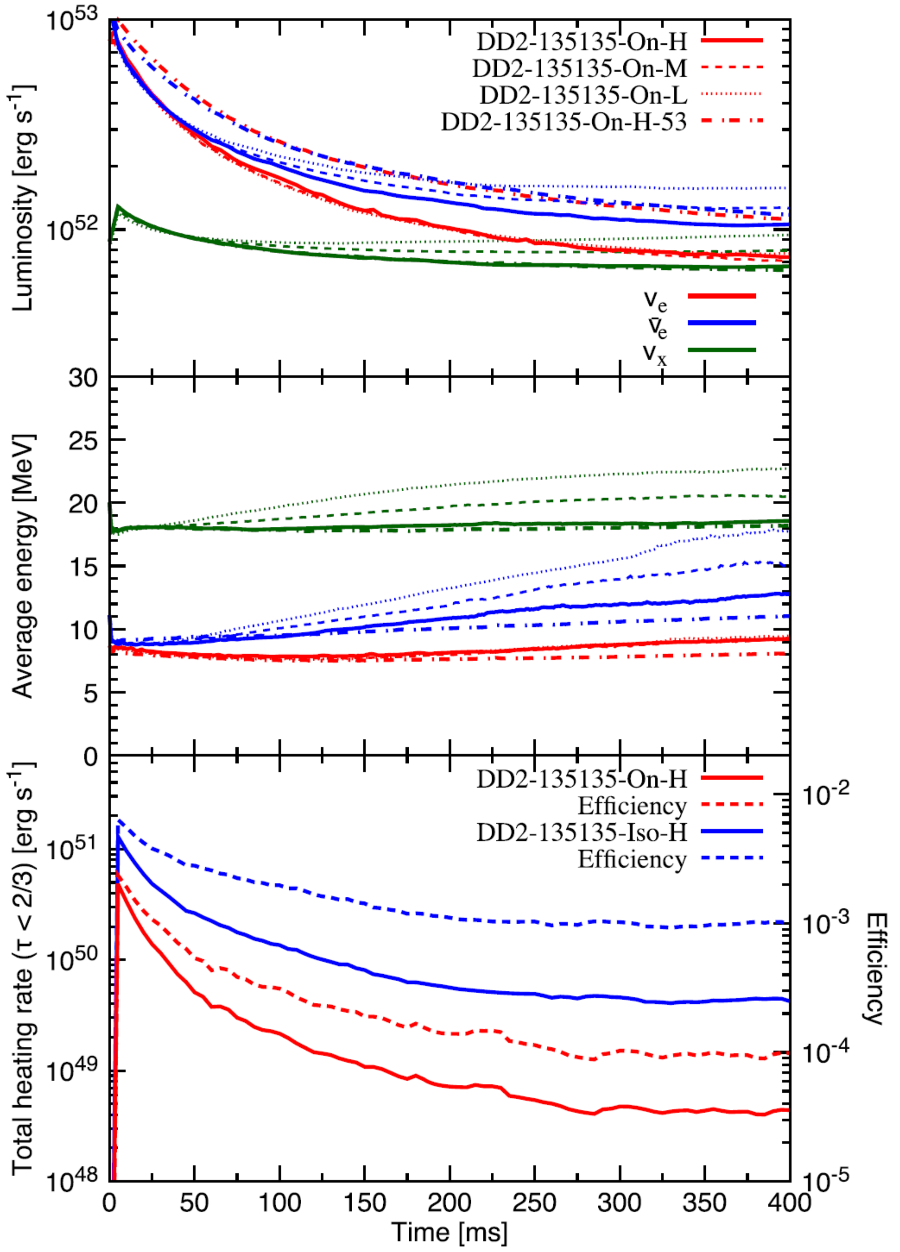


1.35, 1.5, 1.6 M_{sun}

- Neutrino luminosities for three flavors
- Expected detection number for 0.5 Mton HK + KNO
    - > 5 for distance < 5 Mpc in 2-3 sec for hyper-massive NS
    - GW or GRB detectors provide the precise time of the event



- GR axisymmetric neutrino radiation hydrodynamic simulations (Fujibayashi+ 2017)



BNS Merger Rate

- BNS merger rate from aLIGO
  - R_{BNSM} ~1,540 Gpc^{-3}/yr (Abbott+ PRL 2017)
  - For $D_{\mathrm{eff}} = 100\ \mathrm{Mpc}$, $R \cdot \frac{4\pi}{3} D_{\mathrm{eff}}^3$ ~ 6.5 events/yr
  - For $D_{\mathrm{eff}} = 200\ \mathrm{Mpc}$ (detection limit of aLIGO for BNS merger),

$$R \cdot \frac{4\pi}{3} D_{\mathrm{eff}}^3 \sim 51.6 \text{ events/yr}$$

Neutrinos from GRBs

- Detailed models of GRBs still uncertain.
  - How does the central engine produce the fireball and launch the relativistic jet?
- Detection of neutrinos can confirm or reject many scenarios

Detection of Thermal Neutrinos from BNS Merger

- HK + KNO (Kyutoku-Kashiyama 2018)
  - Expected number of neutrino events

$$N_\nu = N_T \int_{t_i}^{t_f} \int_{E_{min}}^{E_{max}} \phi(E,t)\sigma(E)dEdt$$

  - Number of target protons

$$N_T \approx \frac{2}{18}\left(\frac{M_T}{m_p}\right) = 6.7 \times 10^{34}\left(\frac{M_T}{1\ \mathrm{Mt}}\right)$$

  - Cross section for proton

$$\sigma(E) = 9.5 \times 10^{-42}\ \mathrm{cm}\left(\frac{E - 1.3\ \mathrm{MeV}}{10\ \mathrm{MeV}}\right)^2\left(1 - \frac{7E}{m_p c^2}\right)$$

  - Expected number of neutrino events for a single merger

$$n \approx 0.5 \times 10^{-3} \times f_E f_{se} f_{osc}\left(\frac{M_T}{0.5\ \mathrm{Mt}}\right)\left(\frac{E_{\Delta t}}{3 \times 10^{52}\mathrm{erg}}\right)\left(\frac{\langle E\rangle}{10\ \mathrm{MeV}}\right)\left(\frac{D}{100\ \mathrm{Mpc}}\right)^2$$

- Superposition (Kyutoku-Kashiyama 2018)
  - Superposition of many mergers
  - LIGO will tell the time of the events
  - Total number of neutrino detection

$$N = R_{BNSM} \cdot t_{obs} \cdot n$$

$$\sim 0.78 \times \left(\frac{M_T}{0.5\ \mathrm{Mt}}\right)\left(\frac{E_{\Delta t}}{3\times 10^{52}\mathrm{erg}}\right)\left(\frac{\langle E_\nu \rangle}{10\ \mathrm{MeV}}\right)\left(\frac{D_{\mathrm{eff}}}{200\ \mathrm{Mpc}}\right)\left(\frac{t_{obs}}{30\ \mathrm{yr}}\right)$$

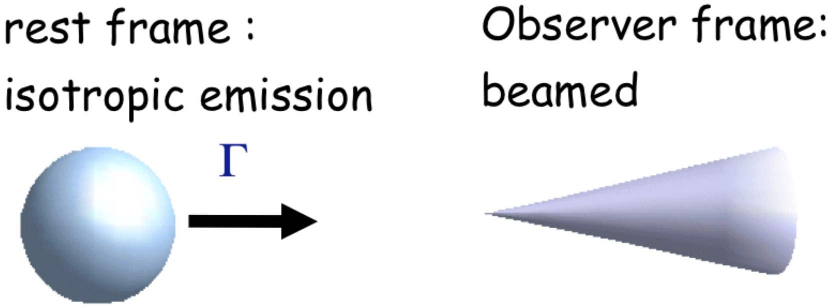
  - Noise control by using timing information from LIGO and using only $\Delta t_{obs} \approx 1$ sec

Relativistic Beaming?

- Particle energy is boosted by the bulk Lorentz factor Γ.

$$\langle E \rangle = \gamma \langle E \rangle_0$$

- Number flux of particles is boosted by Γ^2.
  - Beam is focused into an angle of Γ^{-1}.



$$\frac{E}{D^2} = \Gamma^2 \left(\frac{E}{D^2}\right)_0$$

- Number of events observed decrease by Γ^{-2}.

$$R_{BNSM} = \Gamma^{-2} R^0_{BNSM}$$

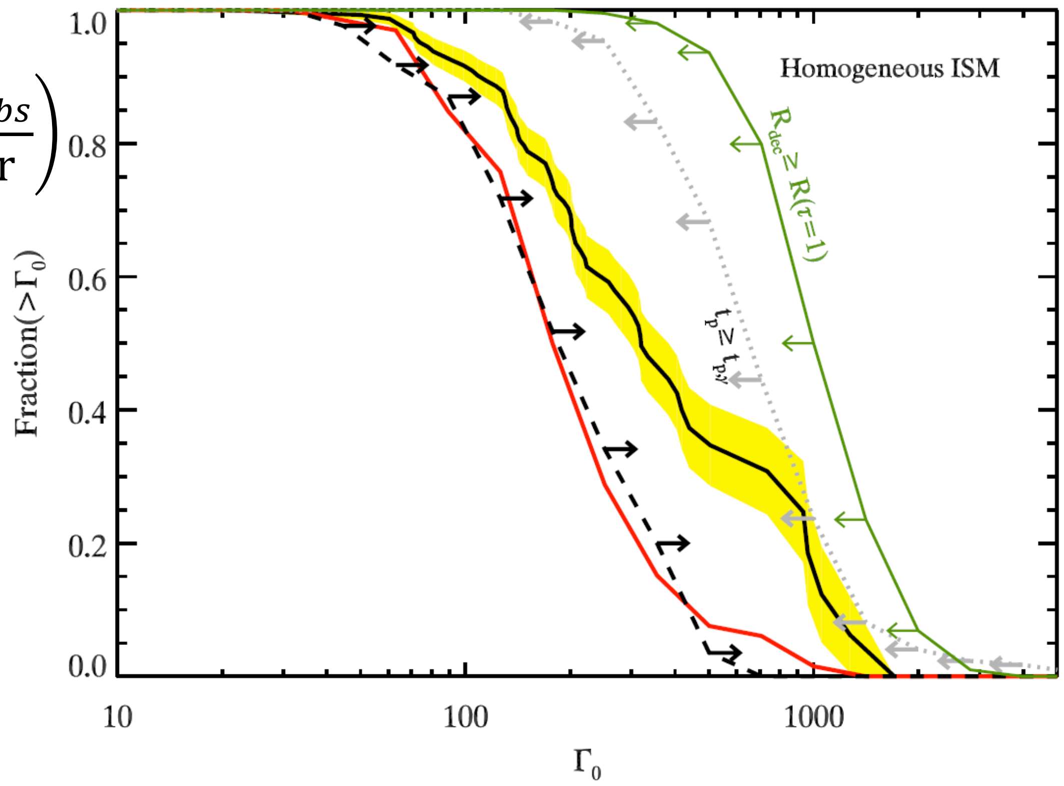
- Duration of burst shortened by Γ^{-1}.

- Total number of neutrino detection

$$N = \Gamma^2 R_{BNSM}^0 \cdot t_{obs} \cdot \Gamma^3 n_0$$

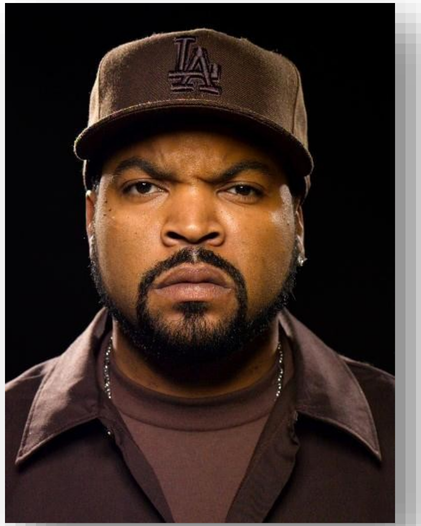
$$\sim 7.8 \left(\frac{\Gamma}{300}\right)\left(\frac{M_T}{0.5\ \mathrm{Mt}}\right)\left(\frac{E_{\Delta t}^0}{3\times10^{52}\mathrm{erg}}\right)\left(\frac{\langle E\rangle_0}{10\ \mathrm{MeV}}\right)\left(\frac{D_{\mathrm{eff}}}{200\ \mathrm{Mpc}}\right)\left(\frac{t_{obs}}{\mathrm{yr}}\right)$$

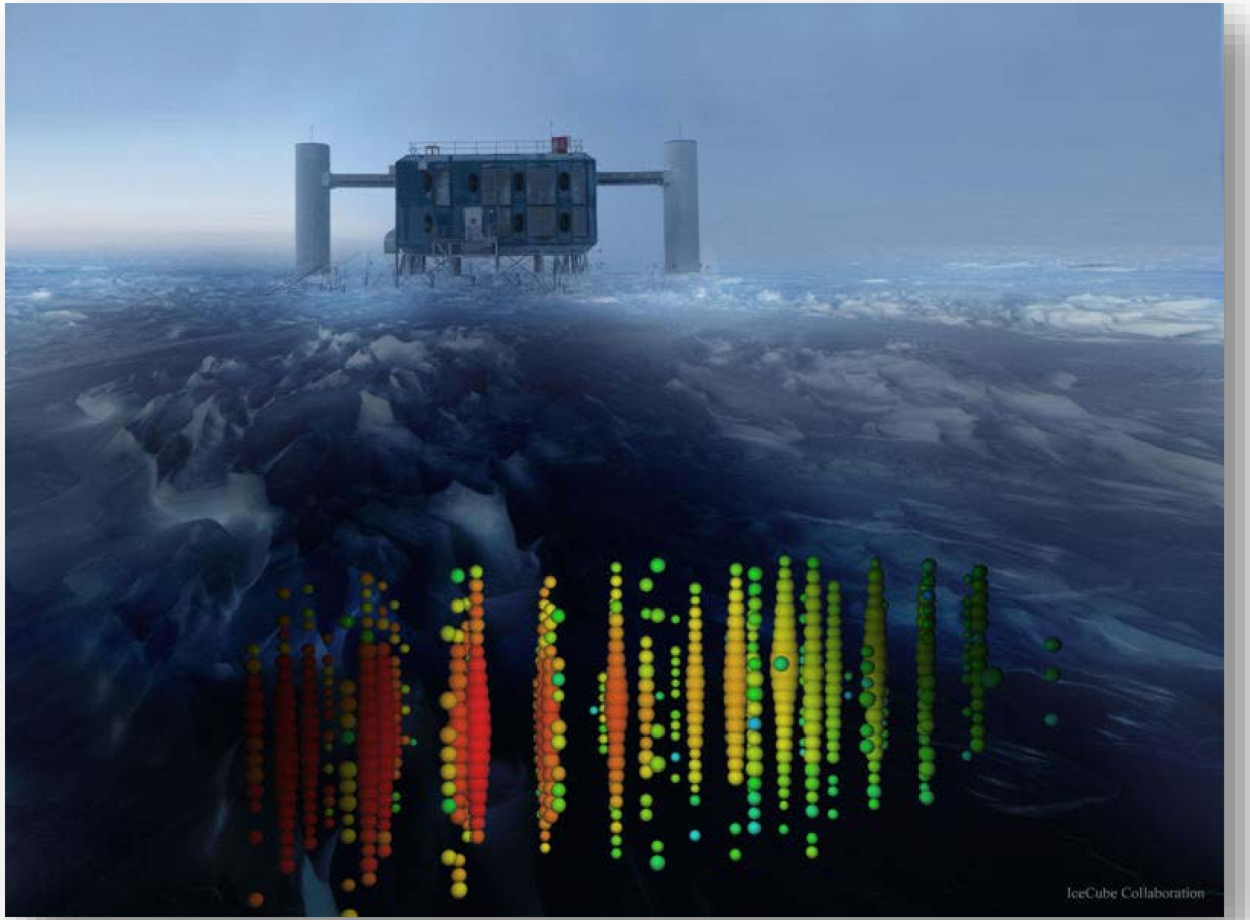
- Estimated Γ
  - $100 \le \Gamma \le 1000$ from observation + afterglow model (Ghirlanda+ 2018)



Neutrinos from Blazar

- IceCube-170922A
  - ~ 290 TeV
  - Consistent with γ-ray blazar TXS 0506+056





RESEARCH ARTICLE

NEUTRINO ASTROPHYSICS

Multimessenger observations of a flaring blazar coincident with high-energy neutrino IceCube-170922A

The IceCube Collaboration, *Fermi*-LAT, MAGIC, *AGILE*, ASAS-SN, HAWC, H.E.S.S., *INTEGRAL*, Kanata, Kiso, Kapteyn, Liverpool Telescope, Subaru, *Swift/NuSTAR*, VERITAS, and VLA/17B-403 teams*†

Previous detections of individual astrophysical sources of neutrinos are limited to the Sun and the supernova 1987A, whereas the origins of the diffuse flux of high-energy cosmic neutrinos remain unidentified. On 22 September 2017, we detected a high-energy neutrino, IceCube-170922A, with an energy of ~290 tera–electron volts. Its arrival direction was consistent with the location of a known γ-ray blazar, TXS 0506+056, observed to be in a flaring state. An extensive multiwavelength campaign followed, ranging from radio frequencies to γ-rays. These observations characterize the variability and energetics of the blazar and include the detection of TXS 0506+056 in very-high-energy γ-rays. This observation of a neutrino in spatial coincidence with a γ-ray–emitting blazar during an active phase suggests that blazars may be a source of high-energy neutrinos.

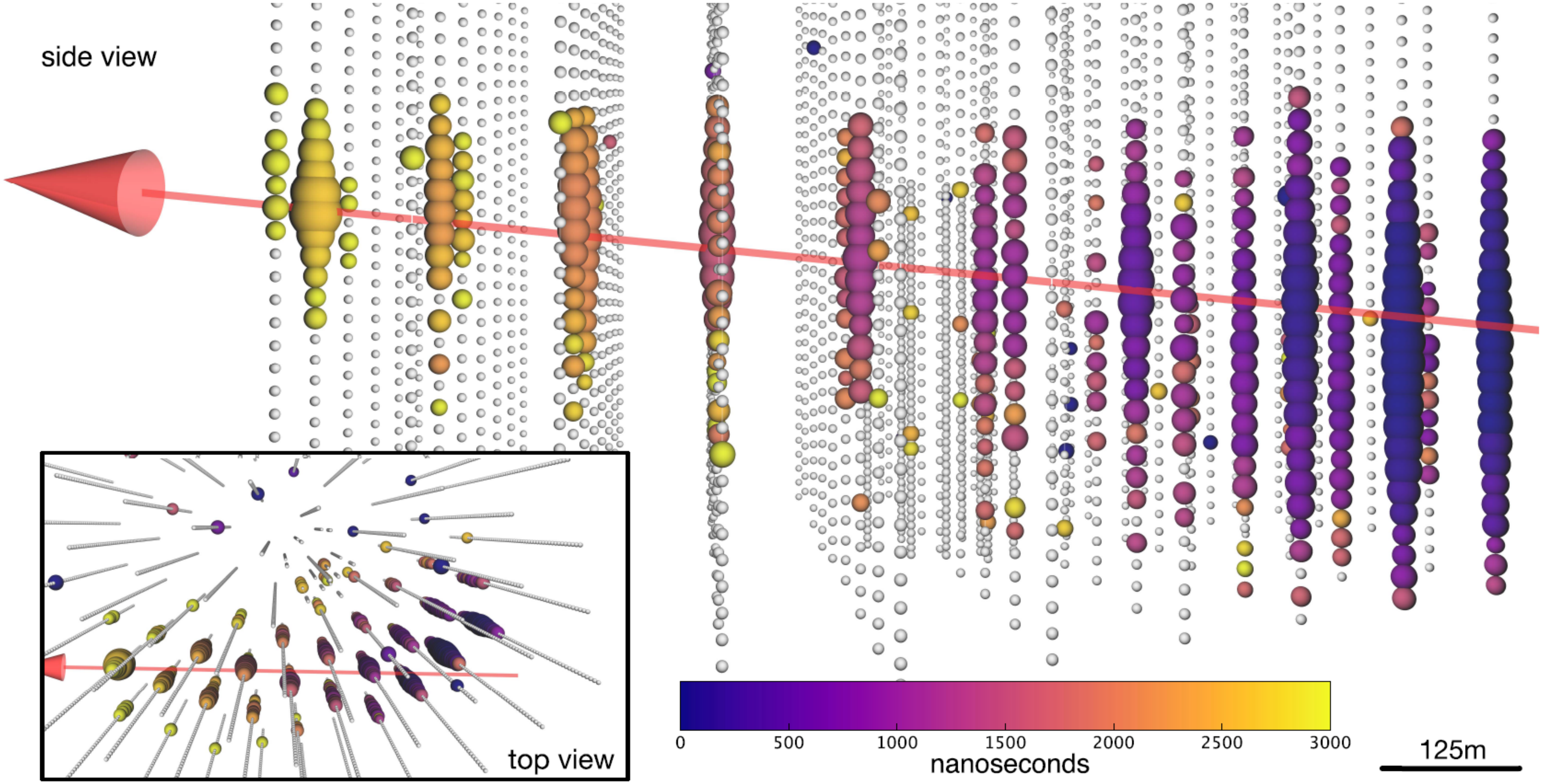
Science, 13 July 2018

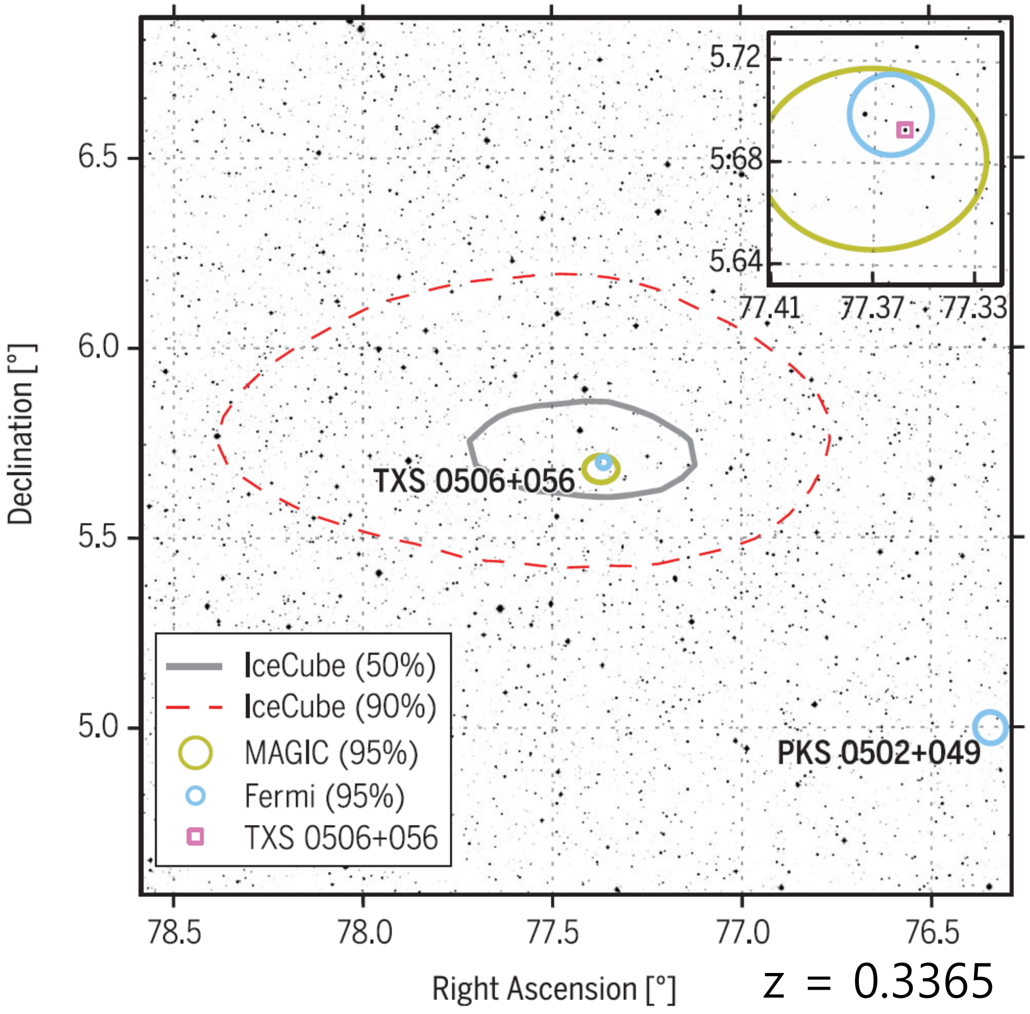
evaluated below, associating neutrino and γ-ray production.

The neutrino alert

IceCube is a neutrino observatory with more than 5000 optical sensors embedded in 1 km^3 of the Antarctic ice-sheet close to the Amundsen-Scott South Pole Station. The detector consists of 86 vertical strings frozen into the ice 125 m apart, each equipped with 60 digital optical modules (DOMs) at depths between 1450 and 2450 m. When a high-energy muon-neutrino interacts with an atomic nucleus in or close to the detector array, a muon is produced moving through the ice at superluminal speed and creating Cherenkov radiation detected by the DOMs. On 22 September 2017 at 20:54:30.43 Coordinated Universal Time (UTC), a high-energy neutrino-induced muon track event was detected in an automated analysis that is part of IceCube's real-time alert system. An automated alert was distributed (*17*) to observers 43 s later, providing an initial estimate of the direction and energy of the event. A sequence of refined reconstruction algorithms was automatically started at the same time, using the full event information. A representation of this neutrino event with the best-fitting reconstructed direction is shown in Fig. 1. Monitoring data from IceCube indicate that the observatory was functioning normally at the time of the event.

R. Morse[32], M. Moulai[14], R. Nahnhauer[52], P. Nakarmi[47], U. Naumann[51], G. Neer[22], H. Niederhausen[45], S. C. Nowicki[23], D. R. Nygren[9], A. Obertacke Pollmann[51], A. Olivas[17], A. O'Murchadha[12], E. O'Sullivan[44], T. Palczewski[9, 8], H. Pandya[37], D. V. Pankova[49], P. Peiffer[33], J. A. Pepper[47], C. Pérez de los Heros[50], D. Pieloth[21], E. Pinat[12], M. Plum[34], P. B. Price[8], G. T. Przybylski[9], C. Raab[12], L. Rädel[1], M. Rameez[20], L. Rauch[52], K. Rawlins[3], I. C. Rea[35], R. Reimann[1], B. Relethford[40], M. Relich[15], E. Resconi[35], W. Rhode[21], M. Richman[40], S. Robertson[2], M. Rongen[1], C. Rott[46], T. Ruhe[21], D. Ryckbosch[26], D. Rysewyk[22], I. Safa[32], T. Sälzer[1], S. E. Sanchez Herrera[23], A. Sandrock[21], J. Sandroos[33], M. Santander[47], S. Sarkar[20, 39], S. Sarkar[23], K. Satalecka[52], P. Schlunder[21], T. Schmidt[17], A. Schneider[32], S. Schoenen[1], S. Schöneberg[11], L. Schumacher[1], S. Sclafani[40], D. Seckel[37], S. Seunarine[42], J. Soedingrekso[21], D. Soldin[37], M. Song[17], G. M. Spiczak[42], C. Spiering[52], J. Stachurska[52], M. Stamatikos[18], T. Stanev[37], A. Stasik[52], R. Stein[52], J. Stettner[1], A. Steuer[33], T. Stezelberger[9], R. G. Stokstad[9], A. Stößl[15], N. L. Strotjohann[52], T. Stuttard[20], G. W. Sullivan[17], M. Sutherland[18], I. Taboada[6], J. Tatar[9, 8], F. Tenholt[11], S. Ter-Antonyan[7], A. Terliuk[52], S. Tilav[37], P. A. Toale[47], M. N. Tobin[32], C. Toennis[46], S. Toscano[13], D. Tosi[32], M. Tselengidou[24], C. F. Tung[6], A. Turcati[35], C. F. Turley[49], B. Ty[32], E. Unger[50], M. Usner[52], J. Vandenbroucke[32], W. Van Driessche[26], D. van Eijk[32], N. van Eijndhoven[13], S. Vanheule[26], J. van Santen[52], E. Vogel[1], M. Vraeghe[26], C. Walck[44], A. Wallace[2], M. Wallraff[1], F. D. Wandler[23], N. Wandkowsky[32], A. Waza[1], C. Weaver[23], M. J. Weiss[49], C. Wendt[32], J. Werthebach[32], S. Westerhoff[32], B. J. Whelan[2], N. Whitehorn[30], K. Wiebe[33], C. H. Wiebusch[1], L. Wille[32], D. R. Williams[47], L. Wills[40], M. Wolf[32], J. Wood[32], T. R. Wood[23], K. Woschnagg[8], D. L. Xu[32], X. W. Xu[7], Y. Xu[45], J. P. Yanez[23], G. Yodh[27], S. Yoshida[15], T. Yuan[32]


side view
top view
0
500
1000
1500
2000
2500
3000
nanoseconds
125m

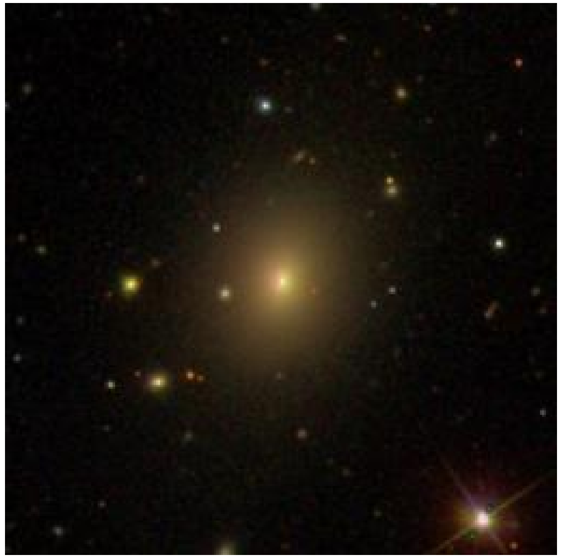


6.5
6.0
5.5
5.0
Declination [°]
78.5
78.0
77.5
77.0
76.5
Right Ascension [°]
5.72
5.68
5.64
77.41
77.37
77.33
TXS 0506+056
PKS 0502+049
IceCube (50%)
IceCube (90%)
MAGIC (95%)
Fermi (95%)
TXS 0506+056

z = 0.3365

Blazar

- Highly variable Active Galactic Nucleus
  - hours to days
- Emission across the EM spectrum
  - Source of high-energy gamma ray
- Some show superluminal motion
- Theoretical model
  - AGN with a relativistic jet, directed very nearly toward Earth
  - Relativistic beaming: $\Gamma \leq 10$
- Classification
  - BL Lac objects: low-power radio galaxies
  - Optically Violently Variable (OVV) Quasars: radio-loud quasars



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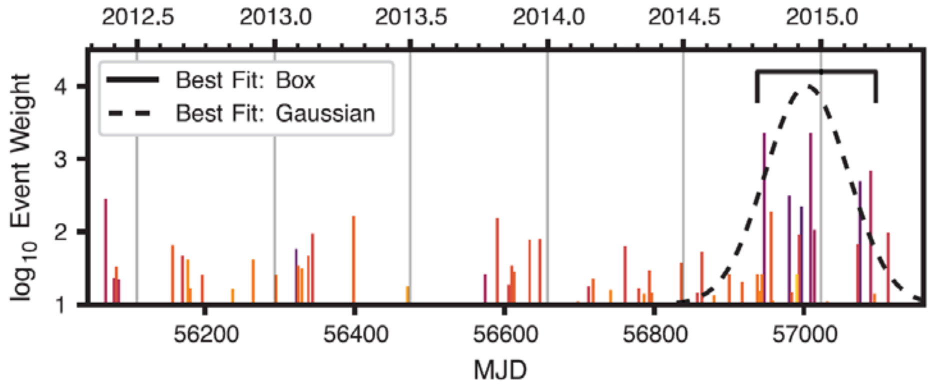
Multi-messenger Observation

- Neutrino luminosity
  - $\langle L_{iso}^{\mu} \rangle \sim 7.2 \times 10^{46}$ erg/s in a 6 month period
  - $E_{iso}^{\mu} \sim 1.1 \times 10^{54}$ erg
- Gamma-ray luminosity

Blazars as neutrino sources

- Neutrinos prior to IceCube-170922A



- All-flavor neutrino fluence: 4.2 x 10^{-3} erg/cm^2 (32 TeV ~ 3.6 PeV)
- E_{iso} = 1.2 x 10^{47} erg/s

- Emission mechanism
  - Accelerated protons interact with ambient lower-energy photons
  - Producing
    - Neutral pions – decay to gamma rays
    - Charged pions – decay to neutrinos and leptons
  - Luminosity and spectrum of lower-energy neutrinos?
    - Assume typical accelerated energy spectrum of protons with zero point provided by IceCube data
    - Calculate the neutrino production rate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