COST Workshop on Higgs and flavour physics: Present and future
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Leptogenesis and the flavour problem

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Why going beyond the SM?

- ❑ Even barring experimental anomalies in $(g-2)_\mu$, B decays,…
- ❑ and in addition to (more or less) compelling theoretical motivations (quantum gravity theory, flavour problem, hierarchy problem, naturalness(?),…)

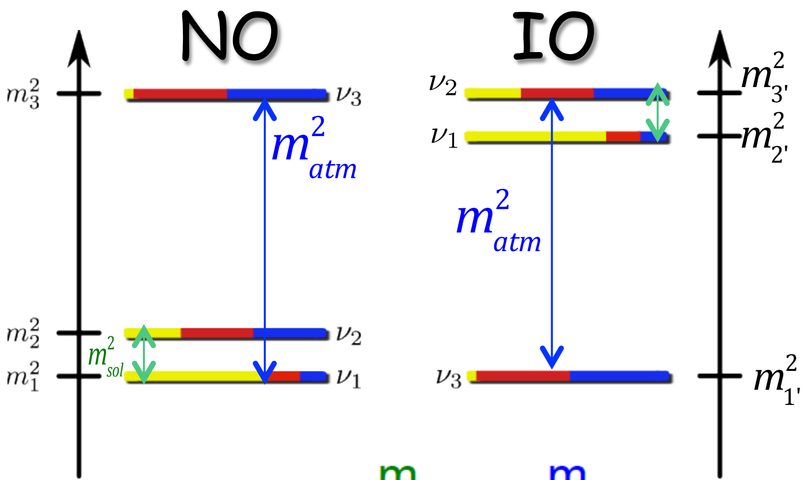
The SM cannot explain:

- *Neutrino masses and mixing*
- *Cosmological Puzzles :*

1. **Dark matter**
2. **Matter - antimatter asymmetry**
3. Inflation
4. Accelerating Universe

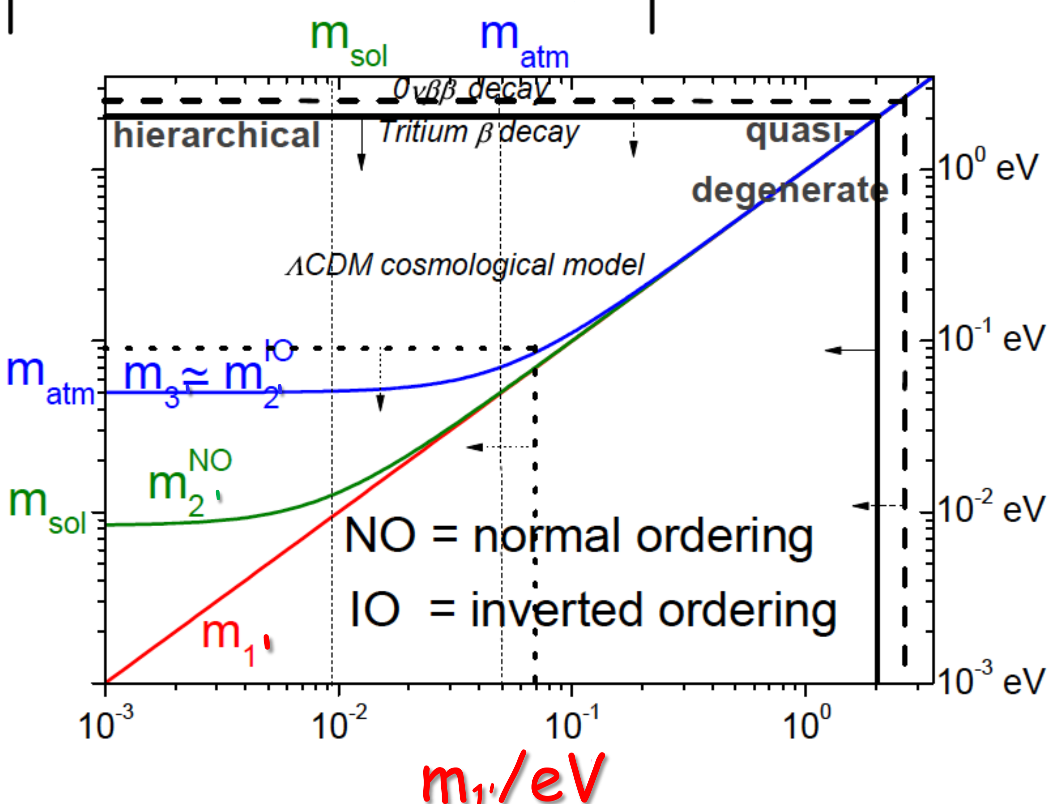
It is reasonable to look for extensions of the SM addressing in a unified picture neutrino masses and mixing and cosmological puzzles and also solve the flavour problem.

Neutrino masses ($m_{1'}<m_{2'}<m_{3'}$)



$$NO: m_2=\sqrt{m_1^2+m_{sol}^2}\,,\quad m_3=\sqrt{m_1^2+m_{atm}^2}$$

$$IO: m_{2'}=\sqrt{m_{1'}^2+m_{atm}^2-m_{sol}^2}\,,\quad m_{3'}=\sqrt{m_{1'}^2+m_{atm}^2}$$



$$m_{sol}=(8.6\pm0.1)\,meV$$

$$m_{atm}=(49.9\pm0.3)\,meV$$

(νfit 2018)

$$\sum_i m_i<0.23\,eV\ (95\%C.L.)$$

(Planck 2015)

$$\Rightarrow m_{1'}\le 0.07\,eV$$

$$\sum_i m_i<0.12\,eV\ (95\%C.L.)$$

(Planck 2018)

$$\Rightarrow m_{1'}\le 0.03\,eV\quad (NO)$$

$$m_{1'}\le 0.016\,eV\quad (IO)$$

Neutrino mixing : $\nu_\alpha = U_{\alpha i}\nu_i$

$$U_{\alpha i} = \begin{pmatrix} U_{e1} & U_{e2} & U_{e3} \\ U_{\mu 1} & U_{\mu 2} & U_{\mu 3} \\ U_{\tau 1} & U_{\tau 2} & U_{\tau 3} \end{pmatrix}$$

Leptonic mixing matrix (PMNS matrix)

$$= \begin{pmatrix} 1 & 0 & 0 \\ 0 & c_{23} & s_{23} \\ 0 & -s_{23} & c_{23} \end{pmatrix} \cdot \begin{pmatrix} c_{13} & 0 & s_{13}e^{-i\delta} \\ 0 & 1 & 0 \\ -s_{13}e^{i\delta} & 0 & c_{13} \end{pmatrix} \cdot \begin{pmatrix} c_{12} & s_{12} & 0 \\ -s_{12} & c_{12} & 0 \\ 0 & 0 & 1 \end{pmatrix} \cdot \begin{pmatrix} e^{i\rho} & 0 & 0 \\ 0 & 1 & 0 \\ 0 & 0 & e^{i\sigma} \end{pmatrix}$$

$\alpha_{31} = 2(\sigma - \rho)$
$\alpha_{21} = -2\rho$

Atmospheric, LB | Reactor, Accel.,LB CP violating phase | Solar,Reactor | bb0ν decay

$c_{ij} = \cos\theta_{ij}$, *and* $s_{ij} = \sin\theta_{ij}$

(νfit November 2018)

NO favoured over IO ($\Delta\chi^2$ (IO-NO)=9.3)

(also talk by M.Tortola)

3σ ranges(NO):

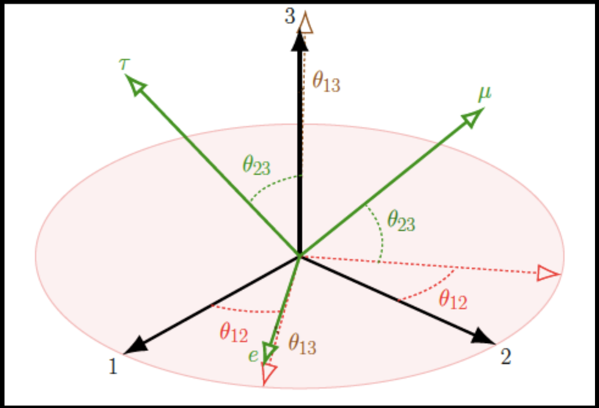
$\theta_{12} = [32^\circ, 36^\circ]$

$\theta_{13} = [8.2^\circ, 9.0^\circ]$

$\theta_{23} = [41^\circ, 52^\circ]$

$\delta = [135^\circ, 366^\circ]$

$\rho, \sigma = [0, 2\pi]$



Minimally extended SM

$$\mathcal{L} = \mathcal{L}_{SM} + \mathcal{L}_Y^{\nu}$$

$$-\mathcal{L}_Y^{\nu} = \overline{\nu_L}\, h^{\nu}\, \nu_R\, \phi \Rightarrow -\mathcal{L}_{\rm mass}^{\nu} = \overline{\nu_L}\, m_D\, \nu_R$$

Dirac mass term

(in a basis where charged lepton mass matrix is diagonal)

diagonalising m_D : $$m_D = V_L^{\dagger} D_{m_D} U_R \qquad D_{m_D} \equiv \begin{pmatrix} m_{D1} & 0 & 0 \\ 0 & m_{D2} & 0 \\ 0 & 0 & m_{D3} \end{pmatrix}$$

$\Rightarrow$ neutrino masses: $m_i = m_{Di}$

leptonic mixing matrix: $U = V_L^{\dagger}$

But many unanswered questions:

- Why neutrinos are much lighter than all other fermions?
- Why large mixing angles (differently from CKM angles)?
- Cosmological puzzles?
- Why not a Majorana mass term as well?

Minimal seesaw mechanism (type I)

• Dirac + (right-right) Majorana mass terms

(Minkowski '77; Gell-mann, Ramond, Slansky; Yanagida; Mohapatra, Senjanovic '79)

$$-\mathcal{L}^{\nu}_{\text{mass}} = \overline{\nu_L}\, m_D\, \nu_R + \frac{1}{2}\overline{\nu_R^c}\, M\, \nu_R + \text{h.c.}$$

violates lepton number

In the see-saw limit ($M \gg m_D$) the mass spectrum splits into 2 sets:

- 3 light **Majorana neutrinos** with masses (seesaw formula):

$$\boxed{\text{diag}(m_1, m_2, m_3) = -U^{\dagger}\, m_D\, \frac{1}{M}\, m_D^T\, U^{\star}}$$

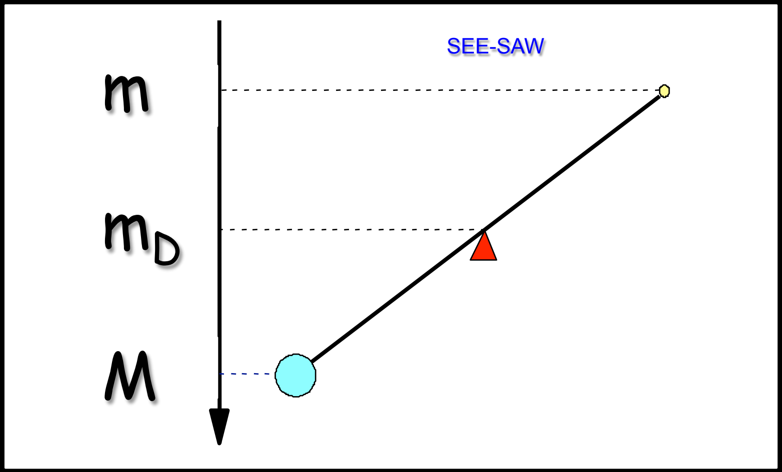
- 3(?) very heavy Majorana neutrinos N_I, N_{II}, N_{III} with $M_{III} > M_{II} > M_I \gg m_D$

1 generation toy model :

$m_D \sim m_{top}$,

$m \sim m_{atm} \sim 50$ meV

$\Rightarrow M \sim M_{GUT} \sim 10^{16}$GeV



Minimal scenario of leptogenesis

(Fukugita, Yanagida '86)

- Type I seesaw mechanism

- Thermal production of RH neutrinos: $T_{RH} \gtrsim T_{lep} \simeq M_i / (2\div 10)$

heavy neutrinos decay $N_i \xrightarrow{\Gamma} L_i + \phi^\dagger \qquad N_i \xrightarrow{\overline{\Gamma}} \overline{L}_i + \phi$

$$\boxed{\text{total CP asymmetries} \quad \varepsilon_i \equiv -\frac{\Gamma - \overline{\Gamma}}{\Gamma + \overline{\Gamma}}} \Rightarrow N_{B-L}^{fin} = \sum_{i=1,2,3} \varepsilon_i \times \kappa_i^{fin}$$

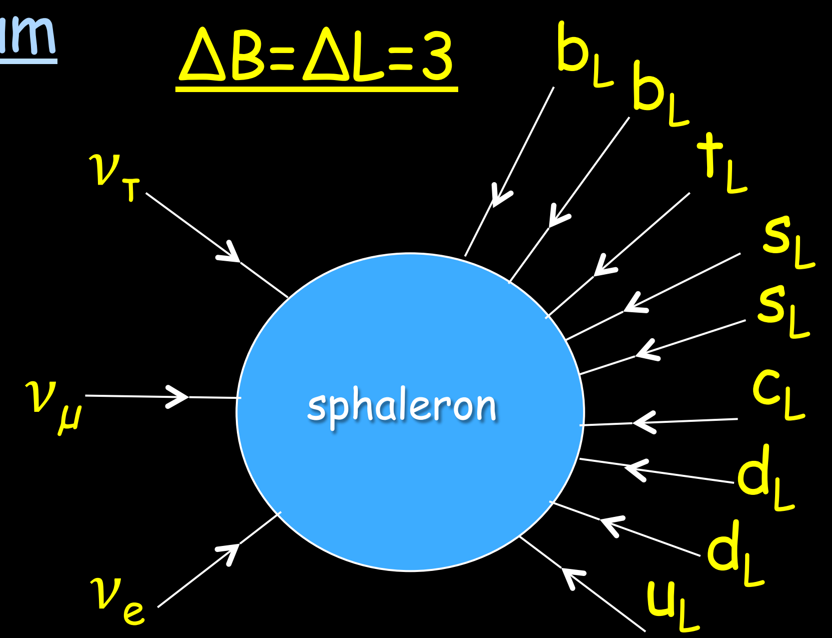
efficiency factors

- Sphaleron processes in equilibrium

$\Rightarrow \; T_{lep} \gtrsim T^{off}_{sphalerons} \sim$ 100 GeV

(Kuzmin, Rubakov, Shaposhnikov '85)

$$\Rightarrow \boxed{\eta_{B0}^{lep} = \frac{a_{sph} N_{B-L}^{fin}}{N_\gamma^{rec}} \simeq 0.01 N_{B-L}^{fin}}$$



Seesaw parameter space

Combining $\eta_{B0}^{lep} \simeq \eta_{B0}^{CMB} \simeq 6\times10^{-10}$ with low energy neutrino data can we test seesaw and leptog.? Problem: too many parameters

(Casas, Ibarra'01) $m_\nu = -m_D \frac{1}{M} m_D^T \Leftrightarrow \boxed{\Omega^T\Omega = I}$ Orthogonal parameterisation

$$\boxed{m_D} = \boxed{U \begin{pmatrix} \sqrt{m_1} & 0 & 0 \\ 0 & \sqrt{m_2} & 0 \\ 0 & 0 & \sqrt{m_3} \end{pmatrix} \Omega \begin{pmatrix} \sqrt{M_1} & 0 & 0 \\ 0 & \sqrt{M_2} & 0 \\ 0 & 0 & \sqrt{M_3} \end{pmatrix}}$$

(in a basis where charged lepton and Majorana mass matrices are diagonal)

light neutrino parameters

heavy neutrino parameters escaping experimental information

- Popular solution: "low-scale" leptogenesis, though no signs so far of new physics at the TeV scale or below able to explain η_{B0}

- High scale leptogenesis is challenging to test but there are a few strategies able to reduce the number of parameters in order to obtain testable predictions on low energy neutrino parameters

Vanilla leptogenesis ⇒ upper bound on ν masses

(Buchmüller, PDB, Plümacher '04; Blanchet, PDB '07)

$$\boxed{\eta_B^{\max}(m_1, M_1) \geq \eta_B^{CMB}}$$

1) Lepton flavor composition is neglected

$$N_i \xrightarrow{\Gamma} \ell_i + \phi^\dagger \qquad N_i \xrightarrow{\bar{\Gamma}} \bar{\ell}_i + \phi$$

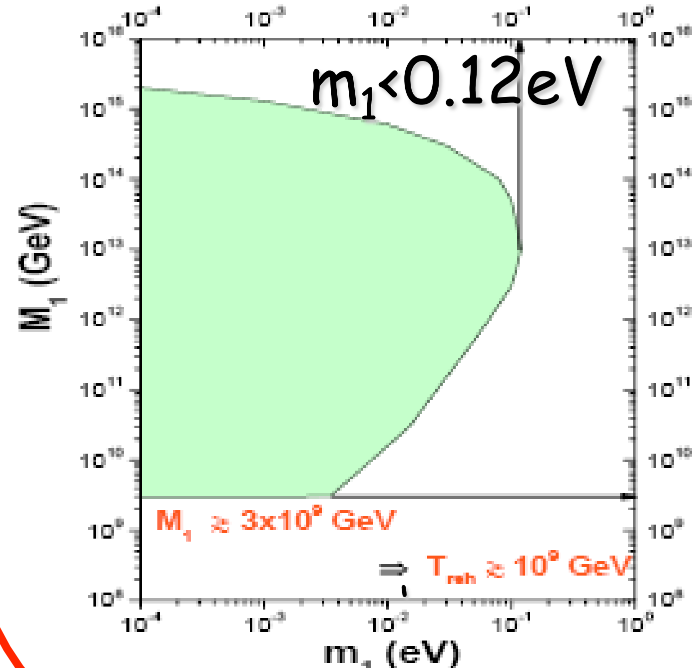
2) Hierarchical spectrum ($M_2 \gtrsim 2M_1$)

3) Strong lightest RH neutrino wash-out

$$\boxed{\eta_{B0} \simeq 0.01 N_{B-L}^{final} \simeq 0.01 \varepsilon_1 \kappa_1^{fin}(K_1, m_1)}$$

decay parameter: $K_1 \equiv \frac{\Gamma_{N_1}(T=0)}{H(T=M_1)}$

All the asymmetry is generated by the lightest RH neutrino



No dependence on the leptonic mixing matrix U: it cancels out

4) Barring fine-tuned cancellations

(Davidson, Ibarra '02)

$$\varepsilon_1 \leq \varepsilon_1^{\max} \simeq 10^{-6} \left(\frac{M_1}{10^{10}\,\text{GeV}}\right) \frac{m_{\text{atm}}}{m_1 + m_3}$$

Independence of the initial conditions (strong thermal leptogenesis)

(Buchmüller, PDB, Plümacher '04)

wash-out of a pre-existing asymmetry $N_{B-L}^{P,\text{initial}}$

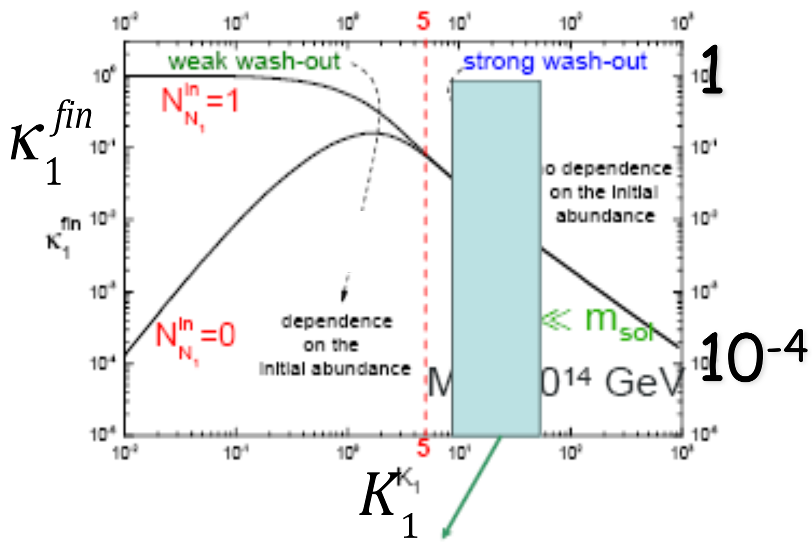
Just a coincidence?

$$\boxed{N_{B-L}^{\rm p,final} = N_{B-L}^{\rm p,initial}\, e^{-\frac{3\pi}{8}\,K_1} \ll N_{B-L}^{\rm f,N_1}}$$

decay parameter: $K_1 \equiv \dfrac{\Gamma_{N_1}}{H(T=M_1)} \sim \dfrac{m_{\rm sol,atm}}{m_\star \sim 10^{-3}\,{\rm eV}} \sim 10 \div 50$

equilibrium neutrino mass: $m_* = \dfrac{16\pi^{5/2}\sqrt{g_*}}{3\sqrt{5}}\dfrac{v^2}{M_{\rm Pl}} \simeq 1.08\times 10^{-3}\ {\rm eV}.$

independence of the initial N_1-abundance as well



$$\boxed{K_{\rm sol} \simeq 9 \lesssim K_1 \lesssim 50 \simeq K_{\rm atm}}$$

Bridging matrix between light and heavy neutrinos

(PDB, M. Re Fiorentin, R. Samanta 1812.07720)

$$N_J \longrightarrow L_J + \phi^\dagger \quad (J = \mathrm{I,II,III}); \qquad |L_J\rangle = B_{iJ}\,|L_i\rangle \quad (i = 1,2,3) \;,$$

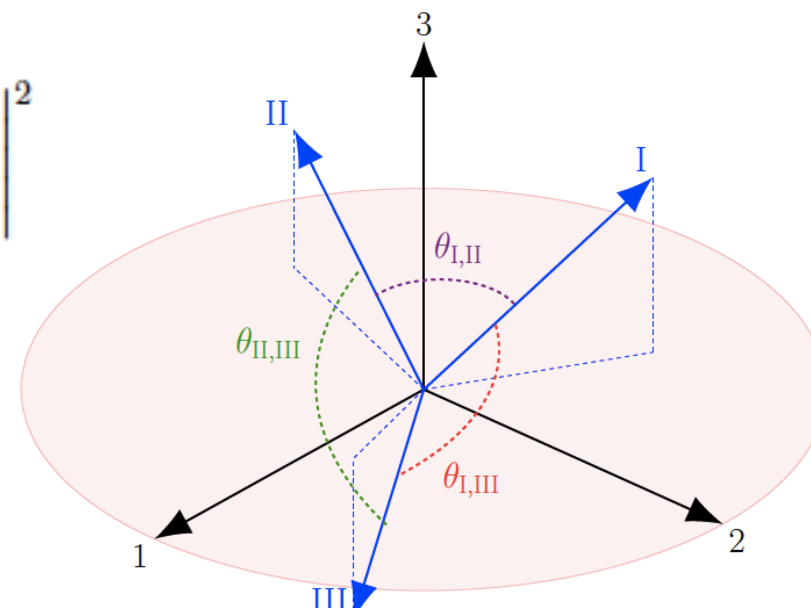
Bridging matrix

$$\boxed{B_{iJ} \equiv \frac{(U^\dagger\, m_D)_{iJ}}{\sqrt{(m_D^\dagger\, m_D)_{JJ}}} = \frac{\sqrt{m_i}\,\Omega_{iJ}}{\sqrt{\sum_k m_k\,|\Omega_{kJ}|^2}}}$$

$$p^0_{J\alpha} \equiv |\langle L_\alpha | L_J\rangle|^2 = \frac{|m_{D\alpha J}|^2}{(m_D^\dagger\, m_D)_{JJ}} = \left|\sum_k U_{\alpha k}\, B_{kJ}\right|^2$$

$$p^0_{IJ} \equiv |\langle L_J | L_I\rangle|^2 = \left|\sum_k B^\star_{kJ}\, B_{kI}\right|^2 = \cos\theta_{IJ}$$

$$If\;\; p^0_{IJ} = \delta_{iJ} \Rightarrow B \;is\; unitary$$



Orthogonal matrix and fine-tuning in seesaw models

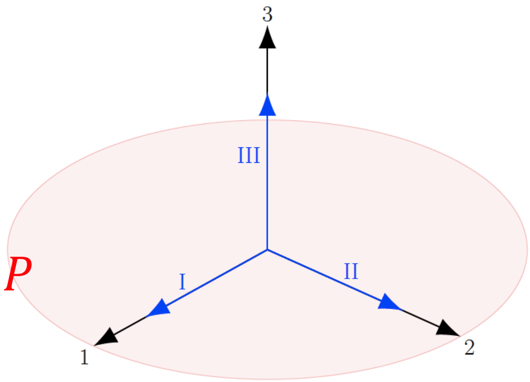
(Lavignac, Masina, Savoy hep-ph/0202086; Branco, Felipe, Joaquim, Masina, Rebelo, hep-ph/0211001; PDB hep-ph/0502082; King, Chen 0903.0125; PDB, Re Fiorentin, Samanta 1812.07720)

In general:

$$m_i = \bar{m}_i \sum_J r_{iJ}\, e^{i\varphi_{iJ}}\,,\quad r_{iJ} = \frac{|\Omega^2_{iJ}|}{\sum_J |\Omega^2_{iJ}|} \propto \frac{1}{M_J}\,,\quad \bar{m}_i = m_i\,\gamma_i\,,\quad \gamma_i = \sum_J |\Omega^2_{iJ}| \geq 1 \text{ (fine-tuning parameters)}$$

Minimum fine-tuning for γ_i=1 for any i ("form dominance"):

$$\Omega = P \Rightarrow m_1 = \frac{m^2_{DI}}{M_I},\ m_2 = \frac{m^2_{DII}}{M_{II}},\ m_3 = \frac{m^2_{DIII}}{M_{III}} + 5\, permutations \Rightarrow B = P$$



In this limit all CP asymmetries (also flavoured) vanish ⇒ NO LEPTOGENESIS

They are typically obtained when a discrete flavour symmetry is imposed ⇒ some symmetry breaking is necessary to have successful leptogenesis

(Jenkins, Manohar 0807.4176; Bertuzzo, PDB, Feruglio, Nardi 0908.0161)

Since SO(3,C) is isomorphic to the restricted Lorentz group $SO^+(3,1)$, any orthogonal matrix can be expressed as a product of a real rotation x boost and the boost is equivalent to a deviation from form dominance:

some "motion" in flavour space is necessary for successful leptogenesis

Charged lepton flavour effects

(Abada et al '06; Nardi et al. '06; Blanchet, PDB, Raffelt '06; Riotto, De Simone '06)

Flavor composition of lepton quantum states matters!

$$|l_1\rangle = \sum_\alpha \langle l_\alpha | l_1\rangle \, |l_\alpha\rangle \qquad (\alpha = e, \mu, \tau)$$

$$|\bar{l}'_1\rangle = \sum_\alpha \langle l_\alpha | \bar{l}'_1\rangle \, |\bar{l}_\alpha\rangle$$

- **T << 10^{12} GeV** ⇒ τ-Yukawa interactions are fast enough break the coherent evolution of $|l_1\rangle$ and $|\bar{l}'_1\rangle$

⇒ incoherent mixture of a τ and of a μ+e components ⇒ 2-flavour regime

- T << 10^9 GeV then also μ-Yukawas in equilibrium ⇒ **3-flavour regime**

M_1

UNFLAVOURED $\qquad N_{B-L}^{final} = \varepsilon_1 \kappa_1^{fin}$

~ 10^{12} GeV — TRANSITION REGIME: DENSITY MATRIX APPROACH NEEDED

2 Flavour regime (τ, e+μ) $\qquad \varepsilon_{1\tau}\kappa_1^{fin}(K_{1\tau}) + \varepsilon_{1e+\mu}\kappa_1^{fin}(K_{1e+\mu})$

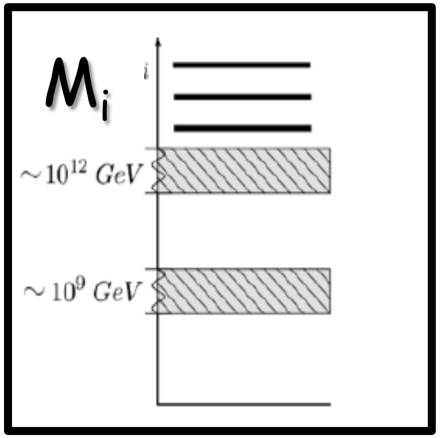
~ 10^9 GeV — TRANSITION REGIME: DENSITY MATRIX APPROACH NEEDED

3 Flavour regime (e, μ, τ) $\qquad \varepsilon_{1\tau}\kappa_1^{fin}(K_{1\tau}) + \varepsilon_{1\mu}\kappa_1^{fin}(K_{1\mu}) + \varepsilon_{1e}\kappa_1^{fin}(K_{1e})$

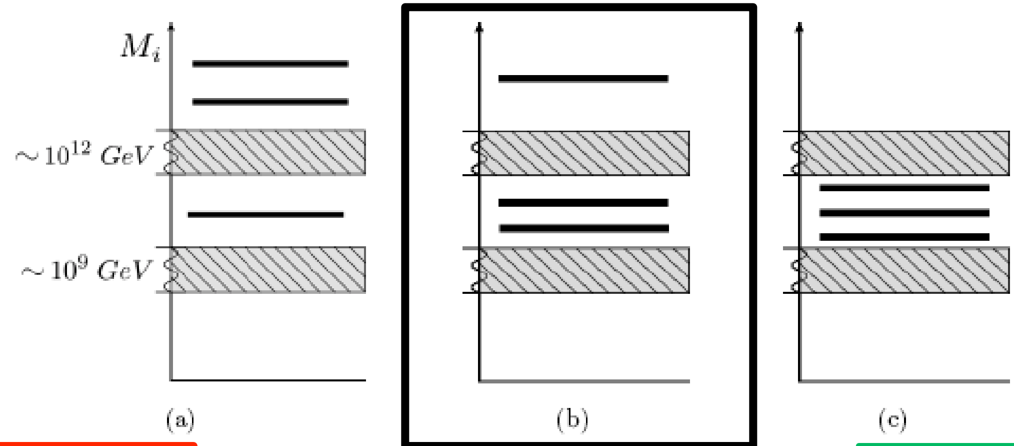
Heavy neutrino lepton flavour effects: 10 scenarios

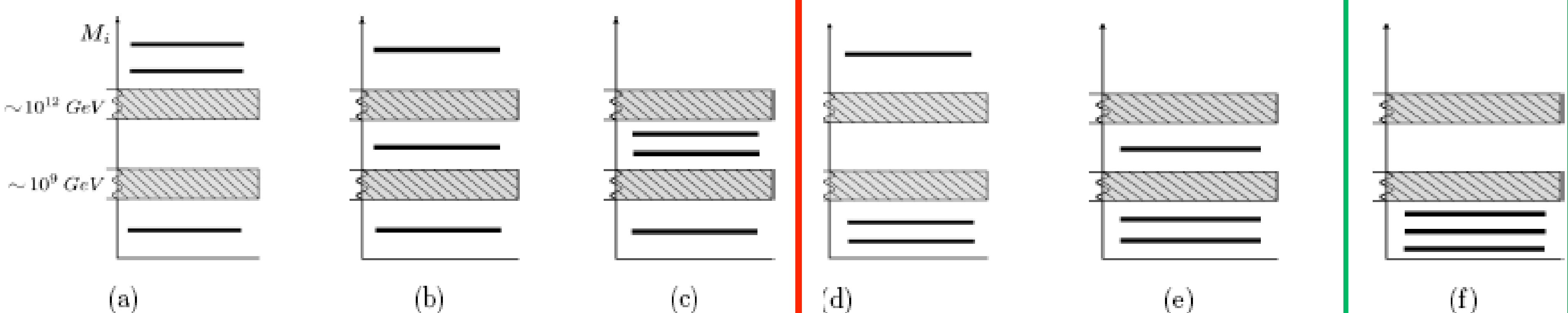
Heavy neutrino flavored scenario

Typically rising in discrete flavour symmetry models



2 RH neutrino scenario





N_2 -dominated scenario:

☛ N_1 produces negligible asymmetry;

Low scale leptogenesis

Example: ARS leptog, (Drewes et al.1711.02862)

An easy limit: all mixing from LH sector

In the flavour basis (both charged lepton mass and Majorana mass matrices are diagonal):

$$-\mathcal{L}^{\nu+\ell}_{\rm mass} = \overline{\alpha_L}\, m_\alpha\, \alpha_R + \overline{\nu_{L\alpha}}\, m_{D\alpha I}\, \nu_{RI} + \frac{1}{2}\,\overline{\nu^c_{RI}}\, M_I\, \nu_{RI} + \text{h.c.}$$

diagonalising again m_D with a bi-unitary transformation: $m_D = V_L^\dagger D_{m_D} U_R$

The seesaw formula becomes:

$$U\, D_m\, U^T = V_L^\dagger\, D_{m_D}\, U_R\, \frac{1}{D_M}\, U_R^T\, D_{m_D}\, V_L^\star$$

$$D_m \equiv diag(m_1, m_2, m_3) \quad D_{m_D} \equiv diag(m_{D1}, m_{D2}, m_{D3}) \quad D_M \equiv diag(M_1, M_2, M_3)$$

<u>AN EASY LIMIT (typically realised imposing a flavour symmetry):</u>

- $U_R = I$ $\Rightarrow$ again $U = V_L^\dagger$ and neutrino masses: $m_i = \dfrac{m_{Di}^2}{M_I}$

If also $m_{D1} = m_{D2} = m_{D3} = \lambda$ then simply: $M_I = \dfrac{\lambda^2}{m_i}$

This limit realises the simple models with $\Omega = P$ (form dominance)

2 RH neutrino models

(King hep-ph/9912492;Frampton,Glashow,Yanagida hep-ph/0208157;Ibarra,Ross2003; Antusch,King,Riotto'08; Antusch, PDB,Jones,King '11; King 1512.07531)

- They can be obtained from 3 RH neutrino models in the limit $M_3 \to \infty$;
- Number of parameters gets reduced to 11;
- Further conditions to get predictions (e.g. $\Omega=P$ +texture zeros)
- Contribution to asymmetry from both 2 RH neutrinos:

the contribution from the lightest (N_1) typically dominates but the contribution from next-to-lightest (N_2) opens new regions that correspons to light sequential dominated neutrino mass models realised in some GUT models. In any case there is still a lower bound

$$M_1 \gtrsim 2\text{x}\,10^{10}\,\text{GeV} \Rightarrow T_{RH} \gtrsim 6 \times 10^9\,\text{GeV}$$

- 2 RH neutrino model realised in A4 x SU(5) SUSY GUT model

with interesting link between "leptogenesis phase" and Dirac phase (F, Bjorkeroth, S.F. King 1505.05504)

- 2 RH neutrino model can be also obtained from 3 RH neutrino models

with 1 vanishing Yukawa eigenvalue ⇒ potential DM candidate

(A.Anisimov, PDB hep-ph/0812.5085)

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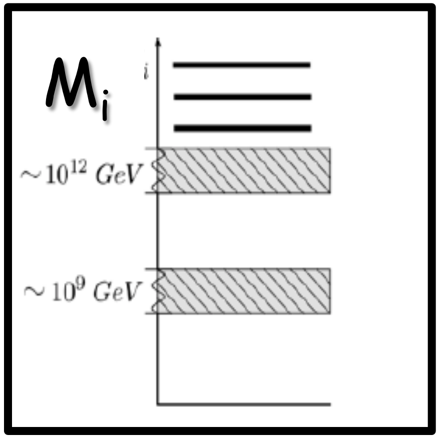
Typically rising in discrete flavour symmetry models

Heavy neutrino flavored scenario



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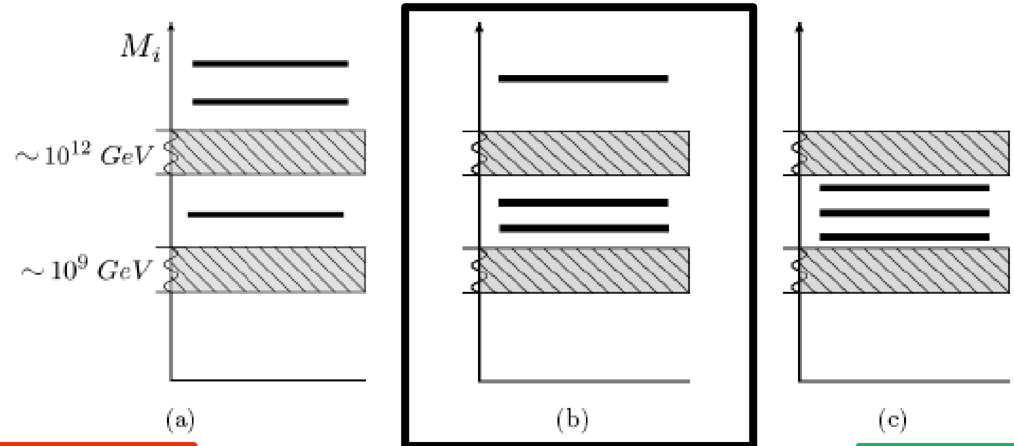

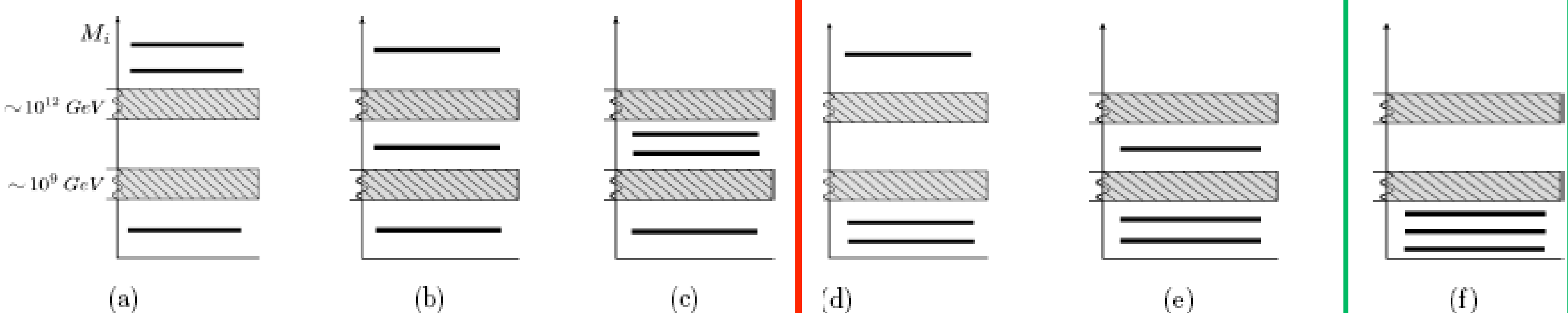

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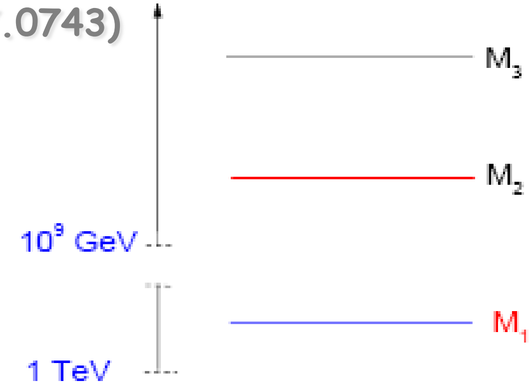
The N_2-dominated scenario

(PDB hep-ph/0502082, Vives hep-ph/0512160;Blanchet,PDB 0807.0743)



- **Unflavoured case:** asymmetry produced from N_2 - RH neutrinos is typically washed-out

$$\eta_{B0}^{lep(N_2)} \simeq 0.01 \cdot \varepsilon_2 \cdot \kappa^{fin}(K_2) \cdot e^{-\frac{3\pi}{8}K_1} << \eta_{B0}^{CMB}$$

- **Adding flavour effects:** lighest RH neutrino wash-out acts on individual flavour ⇒ much weaker

$$\boxed{N_{B-L}^{\rm f}(N_2) = P_{2e}^0\,\varepsilon_2\,\kappa(K_2)\,e^{-\frac{3\pi}{8}\,K_{1e}} + P_{2\mu}^0\,\varepsilon_2\,\kappa(K_2)\,e^{-\frac{3\pi}{8}\,K_{1\mu}} + P_{2\tau}^0\,\varepsilon_2\,\kappa(K_2)\,e^{-\frac{3\pi}{8}\,K_{1\tau}}}$$

- With flavor effects the domain of successful N_2 dominated leptogenesis greatly enlarges: the probability that $K_1 < 1$ is less than 0.1% but the probability that either K_{1e} or $K_{1\mu}$ or $K_{1\tau}$ is less than 1 is ~23%

  (PDB, Michele Re Fiorentin, Rome Samanta)

- Existence of the heaviest RH neutrino N_3 is necessary for the $\varepsilon_{2\alpha}$'s not to be negligible

- It is the only hierarchical scenario that can realise strong thermal leptogenesis (independence of the initial conditions) if the asymmetry is tauon-dominated and if $m_1 \gtrsim 10$ meV (corresponding to $\Sigma_i m_i \gtrsim 80$meV)

  (PDB, Michele Re Fiorentin, Sophie King arXiv 1401.6185)

A less easy limit: SO(10)-inspired models

(Branco et al. '02; Nezri, Orloff '02; Akhmedov, Frigerio, Smirnov '03,PDB, Riotto '08; PDB, Re Fiorentin '12)

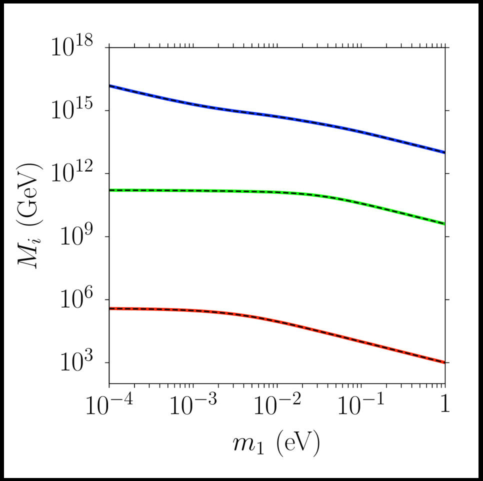
$$U\,D_m\,U^T = V_L^\dagger\,D_{m_D}\,U_R\,\frac{1}{D_M}\,U_R^T\,D_{m_D}\,V_L^\star$$

$$D_m \equiv diag(m_1,m_2,m_3) \quad D_{m_D} \equiv diag(m_{D1},m_{D2},m_{D3}) \quad D_M \equiv diag(M_1,M_2,M_3)$$

- $V_L = I \;\Rightarrow\; M_1 = \dfrac{m_{D1}^2}{m_{\beta\beta}}; \quad M_2 = \dfrac{m_{D2}^2}{m_1 m_2 m_3}\dfrac{m_{\beta\beta}}{|(m_\nu^{-1})_{\tau\tau}|}; \quad M_3 = m_{D3}^2\,|(m_\nu^{-1})_{\tau\tau}|$

If also: $m_{D1} = \alpha_1\, m_{up}; \quad m_{D2} = \alpha_2 m_{charm}; \; m_{D3} = \alpha_3 m_{top}; \quad \alpha_i = O(1)$

Barring fine-tuned solutions, one obtains
a very hierarchical RH neutrino mass spectrum
requiring N_2 -leptogenesis: DOES IT WORK?



The analytical expressions for the M_i 's
can be nicely extended for a generic V_L

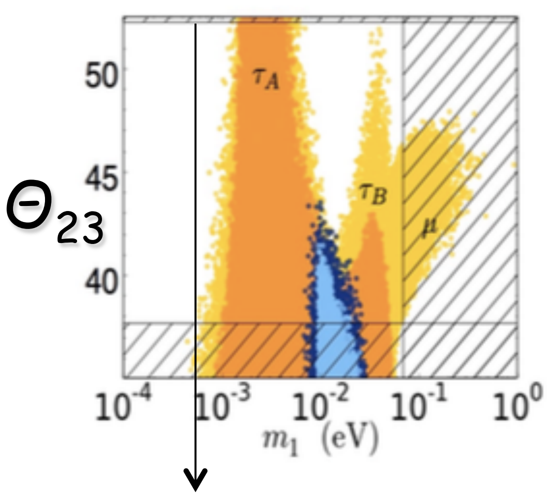
Notice that $U_R = I+\delta$ and for $m_1 \to 0$ one recovers form dominance ($U_R = I$)

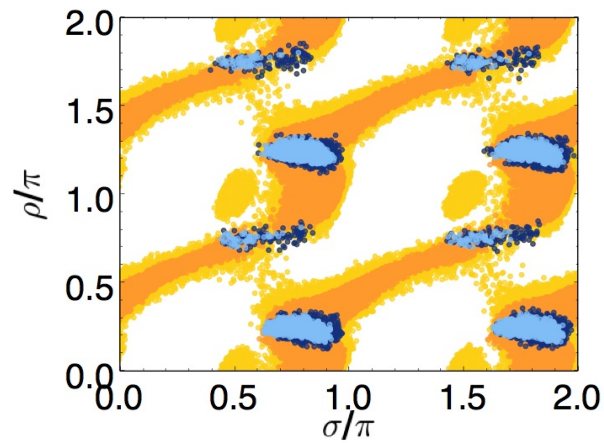
SO(10)-inspired leptogenesis is predictive

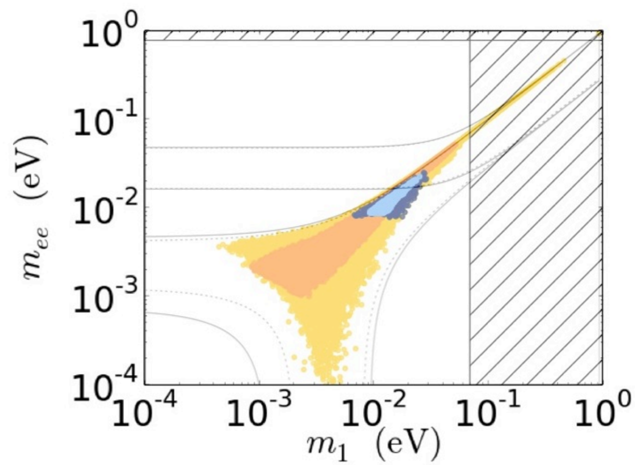
(PDB, Riotto 0809.2285;1012.2343;He,Lew,Volkas 0810.1104)

- dependence on α_1 and α_3 cancels out $\Rightarrow$
  the asymmetry depends only on $\alpha_2 \equiv m_{D2}/m_{charm}$: $\eta_B \propto \alpha_2^2$

$\alpha_2=5$ NORMAL ORDERING $I \leq V_L \leq V_{CKM}$ $V_L = I$







- ➢ Lower bound $m_1 \gtrsim 10^{-3}$ eV
- ➢ Θ_{23} upper bound
- ➢ Majorana phases constrained about specific regions
- ➢ Effective $0\nu\beta\beta$ mass can still vanish but bulk of points above meV

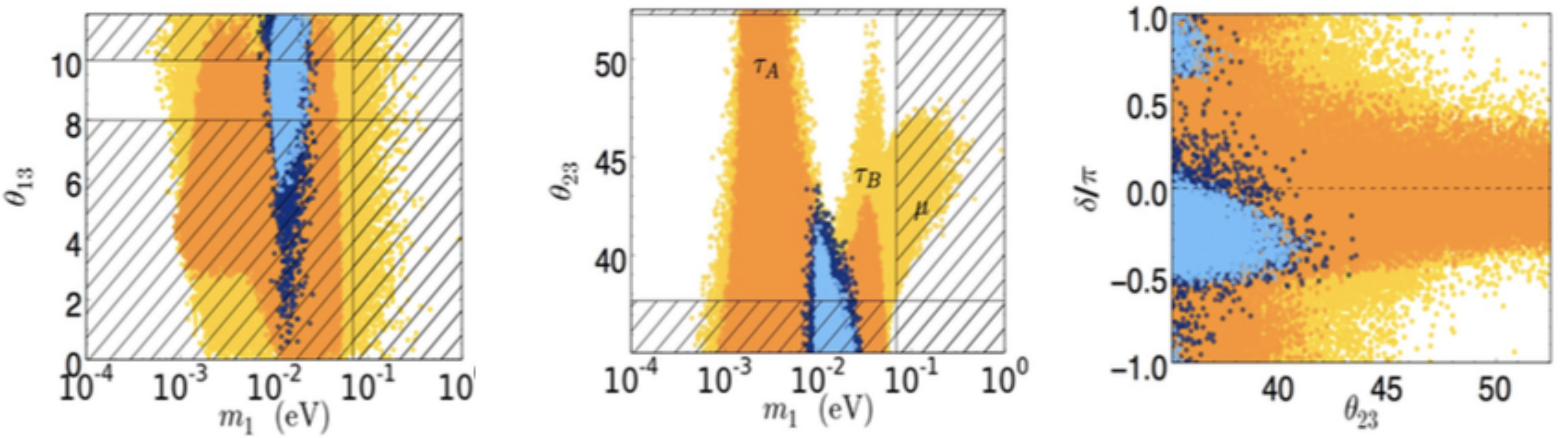
- ➢ only marginal allowed regions for INVERTED ORDERING
- ➢ What are the blue regions?

Strong thermal SO(10)-inspired (STSO10) solution

(PDB,Marzola 09/2011,DESY workshop;1308.1107;PDB,Re Fiorentin,Marzola 1411.5478)

- Strong thermal leptonesis condition can be satisfied for a subset of the solutions only for NORMAL ORDERING

α_2=5 ❑ blue regions: $N_{B-L}^{pre-ex} = 10^{-3}$ (I≤V_L≤V_{CKM}; V_L=I)



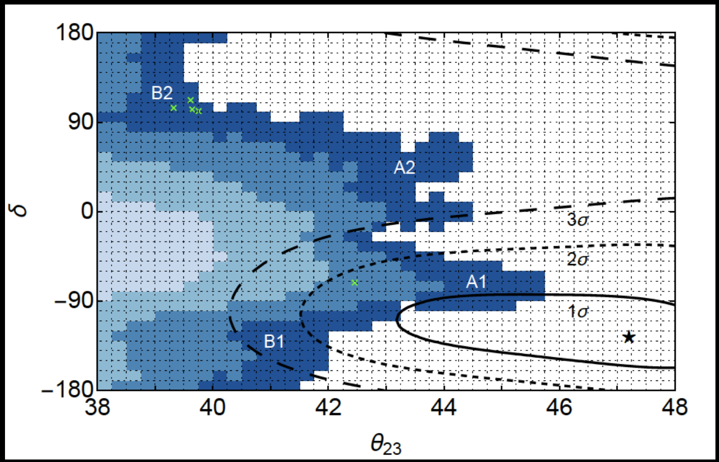
- Absolute neutrino mass scale: $8 \lesssim m_1/\text{meV} \lesssim 30 \Leftrightarrow 70 \lesssim \sum_i m_i/\text{meV} \lesssim 120$
- Non-vanishing Θ_{13};
- Θ_{23} strictly in the first octant;

Strong SO(10)-inspired leptogenesis confronting long baseline experiments

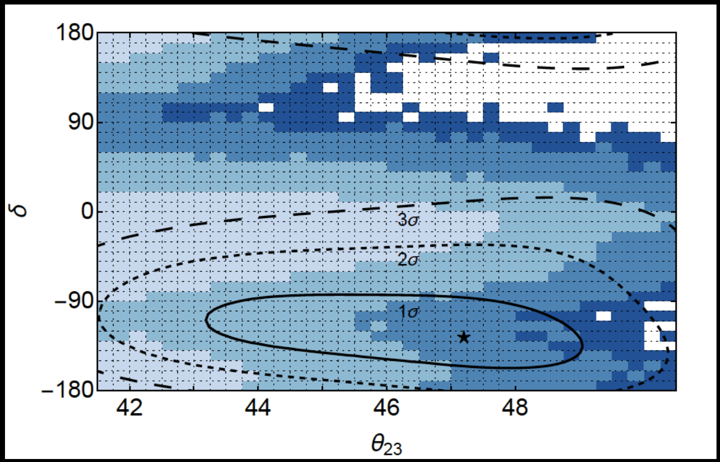
(PDB, Marco Chianese 1802.07690)

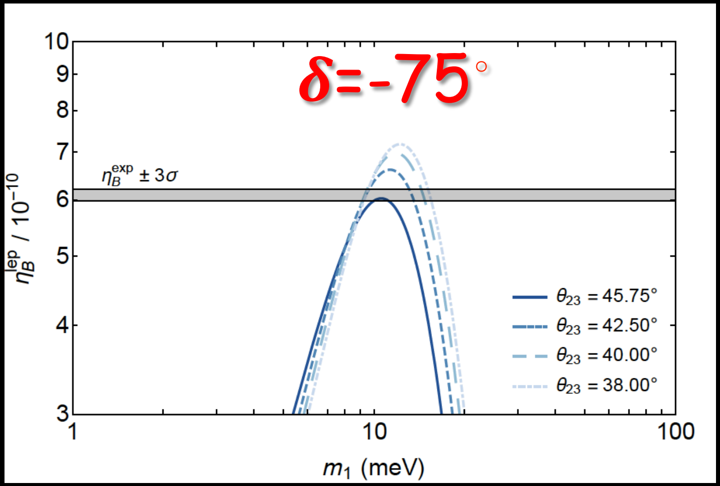
Pre-existing initial asymmetry: $N_{B-L}^{p,i} = 10^{-3}$

$\alpha_2 = m_{D2} / m_{charm} = 5$



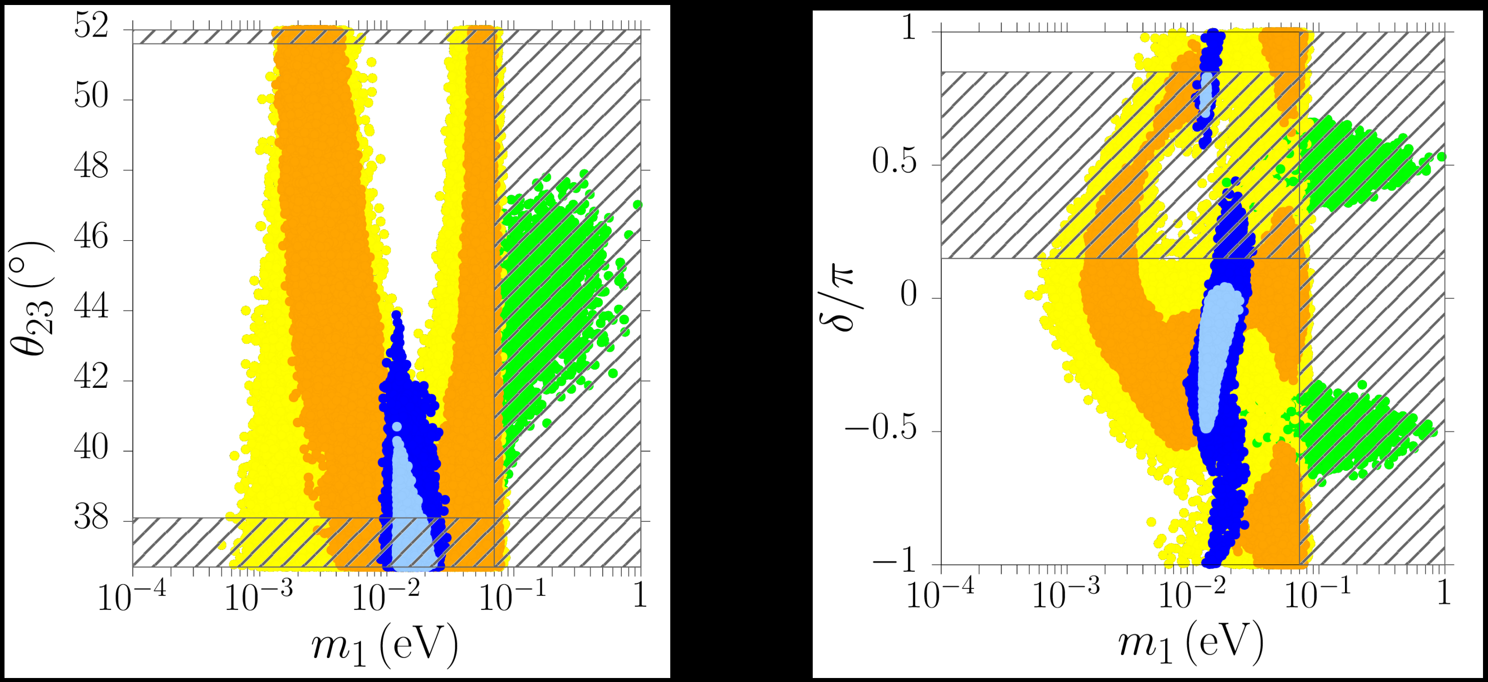
$\alpha_2 = m_{D2} / m_{charm} = 6$





Second octant is compatible with strong thermal condition only if $\alpha_2 \gtrsim 6$: one has to consider realistic models

SO(10)-inspired leptogenesis confronting long baseline and absolute neutrino mass experiments



If the current tendency of data to favour second octant for θ_{23} is confirmed, then SO(10)-inspired leptogenesis predicts a deviation from the hierarchical limit that should be measured at some point (PDB, Samanta in preparation)

A popular class of SO(10) models

(Fritzsch, Minkowski, Annals Phys. 93 (1975) 193-266; R.Slansky, Phys.Rept. 79 (1981) 1-128; G.G. Ross, GUTs, 1985; Dutta, Mimura, Mohapatra, hep-ph/0507319; G. Senjanovic hep-ph/0612312)

In SO(10) models each SM particles generation + 1 RH neutrino are assigned to a single 16-dim representation. Masses of fermions arise from Yukawa interactions of two 16s with vevs of suitable Higgs fields.

The Higgs fields of renormalizable SO(10) models can belong to 10-, 126-,120-dim representations yielding Yukawa part of the Lagrangian

$$\mathcal{L}_Y = 16\,(Y_{10}10_H + Y_{126}\overline{126}_H + Y_{120}120_H)\,16\,.$$

After SSB of the fermions at $M_{GUT}=2\times10^{16}$ GeV one obtains the masses:

up-quark mass matrix $M_u = \boxed{v_{10}^u Y_{10}} + v_{126}^u Y_{126} + v_{120}^u Y_{120}\,,$

down-quark mass matrix $M_d = \boxed{v_{10}^d Y_{10}} + v_{126}^d Y_{126} + v_{120}^d Y_{120}\,,$

neutrino mass matrix $M_D = \boxed{v_{10}^u Y_{10}} - 3v_{126}^u Y_{126} + v_{120}^D Y_{120}\,,$

charged lepton mass matrix $M_l = \boxed{v_{10}^d Y_{10}} - 3v_{126}^d Y_{126} + v_{120}^l Y_{120}\,,$

RH neutrino mass matrix $M_R = v_{126}^R Y_{126}\,,$

LH neutrino mass matrix $M_L = v_{126}^L Y_{126}\,,$

Simplest case but clearly non-realistic: it predicts no mixing at all (both in quark and lepton Sectors). For realistic models one has to add at least the 126 contribution

NOTE: these models do respect SO(10)-inspired conditions

Recent fits within SO(10) models

- Joshipura Patel 2011; Rodejohann, Dueck '13 : **the obtained quite good fits especially including supersymmetry but no leptogenesis and usually compact Spectrum solutions very fine tuned**

- Babu, Bajc, Saad 1612.04329: they find a good fit with NO, hierarchical RH neutrino spectrum but no leptogenesis

- Ohlsson, Pernow 1804.04560: a fit found for NO but minimum χ^2=18.4

- de Anda, King, Perdomo 1710.03229: $SO(10) \times S_4 \times Z_4^R \times Z_4^3$ model: it fits fermion parameters and also find successful leptogenesis respecting the constraints we showed: interesting prediction on neutrinoless double beta decay effective neutrino mass m_{ee} ~11 meV.

- Boucenna, Ohlsson, Pernow 1812.10548: minimal non-supersymmetric SO(10) model with Peccei-Quinn symmetry (best fit χ2=21: tension with θ_{23} in second octant)

In all recent fits a type II term does not seem to help and best fits are type I dominated

An example of realistic model combining GUT+discrete symmetry: SO(10)-inspired leptogenesis in the "A2Z model"

(S.F.King 2014, PDB, S.F.King 1507.06431)

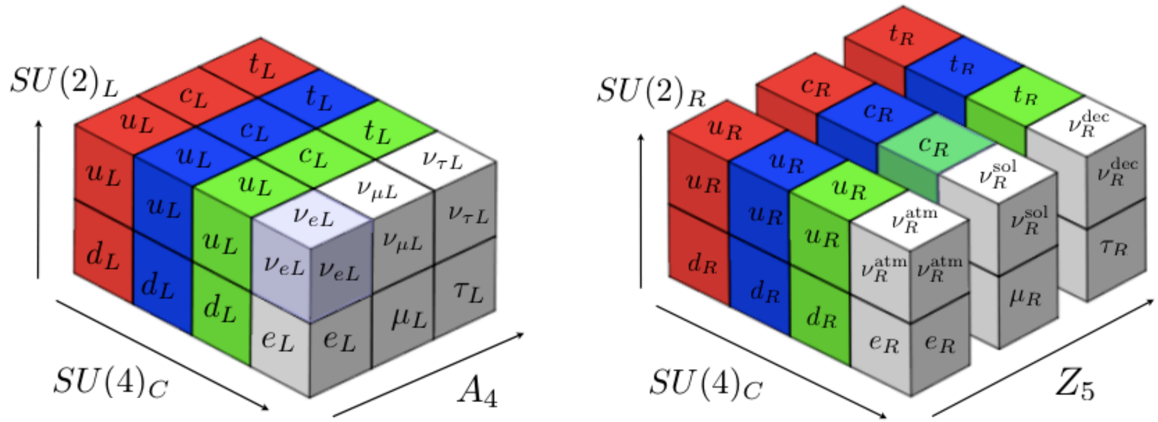
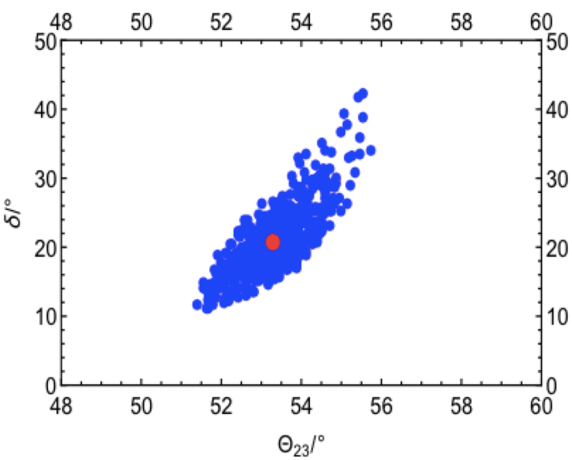
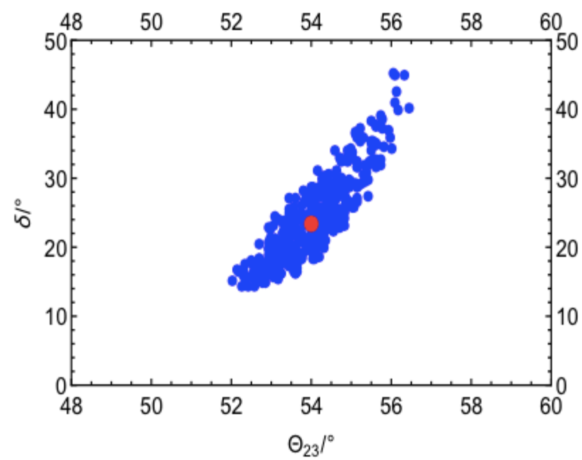


Figure 1: A to Z of flavour with Pati-Salam, where $A \equiv A_4$ and $Z \equiv Z_5$. The left-handed families form a triplet of A_4 and are doublets of $SU(2)_L$. The right-handed families are distinguished by Z_5 and are doublets of $SU(2)_R$. The $SU(4)_C$ unifies the quarks and leptons with leptons as the fourth colour, depicted here as white.





CASE A:

$$m^D_{\nu 1} = m_{\rm up}, \quad m^D_{\nu 2} = m_{\rm charm}, \quad m^D_{\nu 3} = m_{\rm top}$$

CASE B:

$$m^D_{\nu 1} \approx m_{\rm up}, \quad m^D_{\nu 2} \approx 3\, m_{\rm charm}, \quad m^D_{\nu 3} \approx \frac{1}{3}\, m_{\rm top}$$

Conclusion

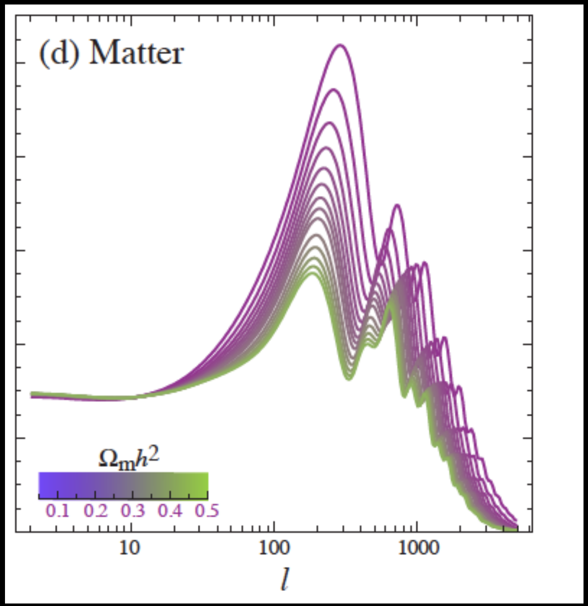
- The quest for an extension of the SM addressing the flavour problem greatly benefits from leptogenesis;
- Low energy leptogenesis can be tested but realistic models able to address the flavour problem typically point to high energy leptogenesis and we need to face the difficulties of testing it: are there realistic models?
- Excluding IO and measuring δ and θ_{23}-45° is of great importance (e.g., strong SO(10)-inspired leptogenesis)
- A detection of non-vanishing absolute neutrino mass scale (not to mention $0\nu\beta\beta$ signal) might prove to be the crucial way to overcome the difficulties, certainly it will test strong thermal leptogenesis that requires $m_1 \gtrsim 10$meV;
- Combining leptogenesis with dark matter provides an alternative strategy in the lack of signals at colliders
- And if you still feel too depressed about it and looking for an inspiration maybe this helps:

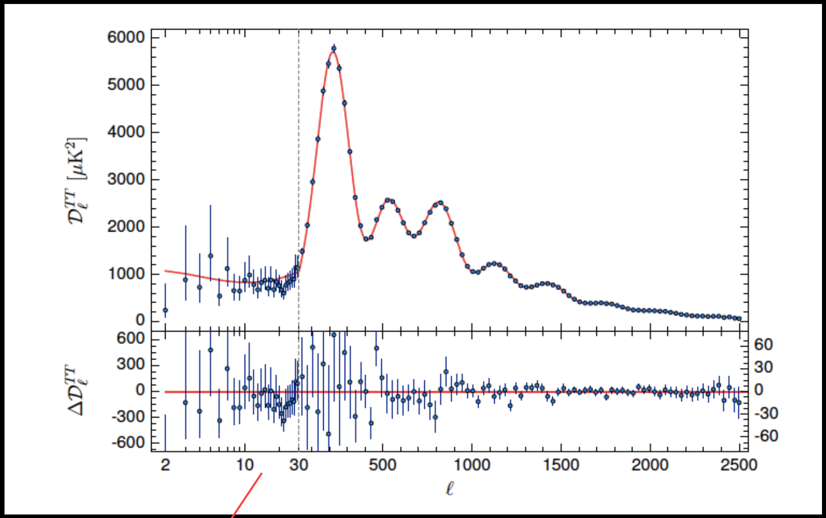
Flavour problem

The Dark Matter of the Universe

(Hu, Dodelson, astro-ph/0110414)

(Planck 2015, 1502.10589)





CMB +"ext"

$$\boxed{\Omega_{CDM,0}h^2 = 0.1188 \pm 0.0010 \sim 5\Omega_{B,0}h^2}$$

A first solution : lowering the scale of the 3 RH neutrinos masses (νMSM)

(Asaka, Blanchet, Shaposhnikov '05)

For $M_1 \ll m_e \Rightarrow$ $\tau_{N_1} = 5 \times 10^{26}\,\mathrm{sec}\left(\frac{M_1}{1\ \mathrm{keV}}\right)^{-5}\left(\frac{\overline{\Theta}^2}{10^{-8}}\right)^{-1}$ >> t_0 $\left(|\overline{\theta}|^2 \equiv \sum_\alpha |m_{D\alpha 1}/M_1|^2\right)$

The production is induced by (non-resonant) RH-LH mixing at T~100 MeV:

$$\Omega_{N_1} h^2 \sim 0.1 \left(\frac{\overline{\theta}}{10^{-4}}\right)^2 \left(\frac{M_1}{keV}\right)^2 \sim \Omega_{DM,0} h^2$$

- The N_1's decay also radiatively and this produces constraints from X-rays (or opportunities to observe it).
- Considering also structure formation constraints, one is forced to consider a resonant production induced by a large lepton asymmetry L ~10^{-4} (3.5 keV line?). (Horiuchi et al. '14; Bulbul at al. '14; Abazajian '14)
- Not clear whether such a large lepton asymmetry can be produced by the same (heavier) RH neutrino decays

An alternative solution: decoupling 1 RH neutrino ⇒ 2 RH neutrino seesaw

(Babu, Eichler, Mohapatra '89; Anisimov, PDB '08)

1 RH neutrino has vanishing Yukawa couplings (enforced by some symmetry such as Z_2):

$$m_D \simeq \begin{pmatrix} 0 & m_{De2} & m_{De3} \\ 0 & m_{D\mu 2} & m_{D\mu 3} \\ 0 & m_{D\tau 2} & m_{D\tau 3} \end{pmatrix}, \text{ or } \begin{pmatrix} m_{De1} & 0 & m_{De3} \\ m_{D\mu 1} & 0 & m_{D\mu 3} \\ m_{D\tau 1} & 0 & m_{D\tau 3} \end{pmatrix}, \text{ or } \begin{pmatrix} m_{De1} & m_{De2} & 0 \\ m_{D\mu 1} & m_{D\mu 2} & 0 \\ m_{D\tau 1} & m_{D\tau 2} & 0 \end{pmatrix},$$

What production mechanism? Turning on tiny Yukawa couplings?

Yukawa basis:

$$m_D = V_L^\dagger \, D_{m_D} \, U_R \, .$$

$$D_{m_D} \equiv v \, \mathrm{diag}(h_A, h_B, h_C), \text{ with } h_A \le h_B \le h_C .$$

$$\tau_{\rm DM} = \frac{4\pi}{h_A^2 \, M_{\rm DM}} \simeq 0.87 \, h_A^{-2} \, 10^{-23} \left(\frac{\rm GeV}{M_{\rm DM}}\right) \rm s \quad \Rightarrow \quad \boxed{\tau_{DM} > \tau_{DM}^{\min} \simeq 10^{28}\, s \Rightarrow h_A < 3\times 10^{-26} \sqrt{\frac{GeV}{M_{DM}} \times \frac{10^{28}\, s}{\tau_{DM}^{\min}}}}$$

One could think of an abundance induced by RH neutrino mixing, considering that:

$$N_{DM} \simeq 10^{-9} (\Omega_{DM,0} h^2) N_\gamma^{prod} \frac{TeV}{M_{DM}}$$

It would be enough to convert just a tiny fraction of ("source") thermalised RH neutrinos but it still does not work with standard Yukawa couplings

Proposed production mechanisms

Starting from a 2 RH neutrino seesaw model

$$m_D \simeq \begin{pmatrix} 0 & m_{De2} & m_{De3} \\ 0 & m_{D\mu 2} & m_{D\mu 3} \\ 0 & m_{D\tau 2} & m_{D\tau 3} \end{pmatrix}, \text{ or } \begin{pmatrix} m_{De1} & 0 & m_{De3} \\ m_{D\mu 1} & 0 & m_{D\mu 3} \\ m_{D\tau 1} & 0 & m_{D\tau 3} \end{pmatrix}, \text{ or } \begin{pmatrix} m_{De1} & m_{De2} & 0 \\ m_{D\mu 1} & m_{D\mu 2} & 0 \\ m_{D\tau 1} & m_{D\tau 2} & 0 \end{pmatrix},$$

many production mechanisms have been proposed:

- from $SU(2)_R$ extra-gauge interactions (LRSM) ;
- from inflaton decays (Anisimov,PDB'08; Higaki, Kitano, Sato '14);
- from resonant annihilations through SU(2)' extra-gauge interactions (Dev, Kazanas,Mohapatra,Teplitz, Zhang '16);
- From new $U(1)_Y$ interactions connecting DM to SM (Dev, Mohapatra,Zhang '16);
- From $U(1)_{B-L}$ interactions (Okada, Orikasa '12);
- …………………

In all these models IceCube data are fitted through fine tuning of parameters responsible for decays (they are post-dictive)

RH neutrino mixing from Higgs portal

(Anisimov,PDB '08)

Assume new interactions with the standard Higgs:

$$\mathcal{L} = \frac{\lambda_{IJ}}{\Lambda}\,\phi^\dagger\,\phi\,\overline{N_I^c}\,N_J \qquad \text{(I,J=A,B,C)}$$

In general they are non-diagonal in the Yukawa basis: this generates a RH neutrino mixing. Consider a 2 RH neutrino mixing for simplicity and consider medium effects:

From the Yukawa interactions: $V_J^Y = \frac{T^2}{8\,E_J}\,h_J^2$

From the new interactions: $V_{JK}^\Lambda \simeq \frac{T^2}{12\,\Lambda}\,\lambda_{JK}$

effective mixing Hamiltonian (in monocromatic approximation)

$$\Delta H \simeq \begin{pmatrix} -\frac{\Delta M^2}{4p} - \frac{T^2}{16p}\,h_S^2 & \frac{T^2}{12\Lambda} \\ \frac{T^2}{12\Lambda} & \frac{\Delta M^2}{4p} + \frac{T^2}{16p}\,h_S^2 \end{pmatrix} \Rightarrow \sin 2\theta_\Lambda^{\rm m} = \frac{\sin 2\theta_\Lambda}{\sqrt{(1+v_S^Y)^2 + \sin^2 2\theta_\Lambda}}$$

$$\Delta M^2 \equiv M_S^2 - M_{\rm DM}^2.$$

$$v_S^Y \equiv T^2\,h_S^2/(4\,\Delta M^2)$$

If $\Delta m^2 < 0$ ($M_{DM} > M_S$) there is a resonance for $v_S^Y=-1$ at:

$$z_{\rm res} \equiv \frac{M_{\rm DM}}{T_{\rm res}} = \frac{h_S\,M_{\rm DM}}{2\sqrt{M_{\rm DM}^2 - M_S^2}}$$

Non-adiabatic conversion

(Anisimov,PDB '08; P.Ludl.PDB,S.Palomarez-Ruiz '16)

Adiabaticity parameter at the resonance

$$\gamma_{\rm res} \equiv \left. \frac{|E^{\rm m}_{\rm DM} - E^{\rm m}_{\rm S}|}{2\,|\dot{\theta}_m|} \right|_{\rm res} = \sin^2 2\theta_\Lambda(T_{\rm res})\, \frac{|\Delta M^2|}{12\, T_{\rm res}\, H_{\rm res}} \,,$$

Landau-Zener formula
(more accurate calculation
employing density matrix
Solution is needed
PDB,Farrag,Katori in prep)

$$\left. \frac{N_{N_{\rm DM}}}{N_{N_{\rm S}}} \right|_{\rm res} \simeq \frac{\pi}{2}\, \gamma_{\rm res}$$

(remember that we need only a small fraction to be converted so necessarily γ_{res}<<<1)

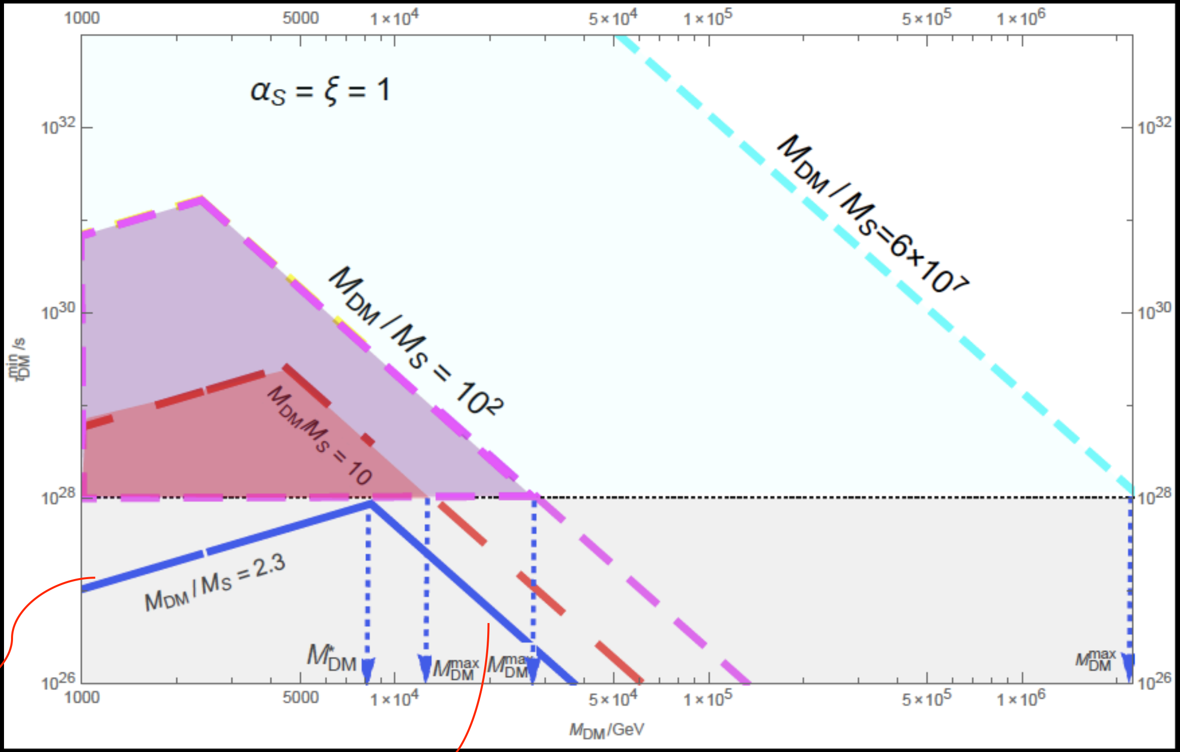
$$\Rightarrow \quad \Omega_{\rm DM}\, h^2 \simeq \frac{0.15}{\alpha_{\rm S}\, z_{\rm res}} \left(\frac{M_{\rm DM}}{M_{\rm S}}\right) \left(\frac{10^{20}\,{\rm GeV}}{\widetilde{\Lambda}}\right)^2 \left(\frac{M_{\rm DM}}{\rm GeV}\right)$$

For successful dark-matter genesis $\Rightarrow$

$$\widetilde{\Lambda}_{\rm DM} \simeq 10^{20} \sqrt{\frac{1.5}{\alpha_{\rm S}\, z_{\rm res}}\, \frac{M_{\rm DM}}{M_{\rm S}}\, \frac{M_{\rm DM}}{\rm GeV}}\ {\rm GeV}$$

2 options: either Λ<M_{Pl} and λ_{AS}<<< 1 or λ_{AS}~ 1 and Λ>>>M_{Pl} :
it is possible to think of models in both cases.

Decays: a natural allowed window on M_{DM}



Lower bound from 2 body decays

Upper bound from 4 body decays

Increasing M_{DM}/M_S relaxes the constraints since it allows higher T_{res} (⇒more efficient production) keeping small N_S Yukawa coupling (helping stability)! But there Is an upper limit to T_{res} from usual upper limit on reheat temperature.

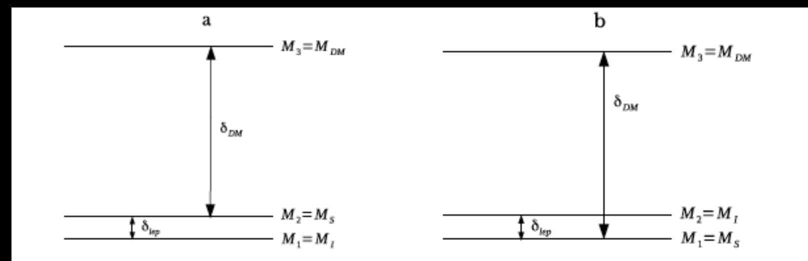
Unifying Leptogenesis and Dark Matter

(PDB, NOW 2006;Anisimov,PDB,0812.5085;PDB, P.Ludl,S. Palomarez-Ruiz 1606.06238+see recent v3)

- Interference between N_A and N_B can give sizeable CP decaying asymmetries able to produce a matter-antimatter asymmetry but since $M_{DM}>M_S$ necessarily $N_{DM}=N_3$ and $M_1 \simeq M_2$ ⇒ leptogenesis with quasi-degenerate neutrino masses

$$\delta_{DM} \equiv (M_3 - M_S)/M_S$$

$$\delta_{lep} \equiv (M_2 - M_1)/M_1$$



$$\varepsilon_{i\alpha} \simeq \frac{\overline{\varepsilon}(M_i)}{K_i}\left\{\mathcal{I}_{ij}^{\alpha}\,\xi(M_j^2/M_i^2) + \mathcal{J}_{ij}^{\alpha}\,\frac{2}{3(1-M_i^2/M_j^2)}\right\}$$

(Covi,Roulet,Visssani '96)

$$\overline{\varepsilon}(M_i) \equiv \frac{3}{16\,\pi}\left(\frac{M_i\,m_{\rm atm}}{v^2}\right) \simeq 1.0\times 10^{-6}\left(\frac{M_i}{10^{10}\,{\rm GeV}}\right),$$

$$\xi(x) = \frac{2}{3}x\left[(1+x)\ln\left(\frac{1+x}{x}\right) - \frac{2-x}{1-x}\right],$$

Analytical expression for the asymmetry:

$$\eta_B \simeq 0.01\,\frac{\overline{\varepsilon}(M_1)}{\delta_{\rm lep}}\,f(m_\nu,\Omega)\,,$$

Efficiency factor

$$f(m_\nu,\Omega) \equiv \frac{1}{3}\left(\frac{1}{K_1}+\frac{1}{K_2}\right)\sum_\alpha \kappa(K_{1\alpha}+K_{2\alpha})\,[\mathcal{I}_{12}^{\alpha}+\mathcal{J}_{12}^{\alpha}]\,,$$

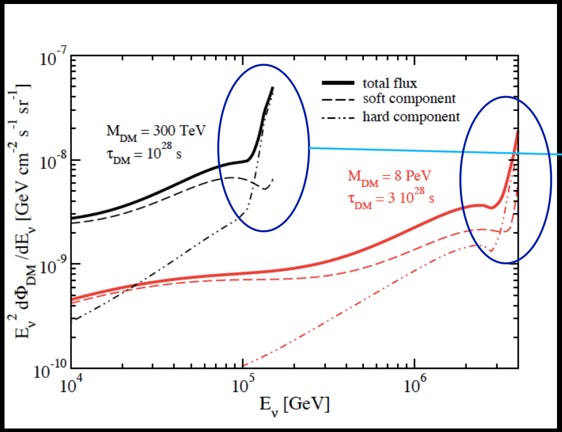
- $M_S \gtrsim 2\, T_{sph} \simeq 300$ GeV ⇒ $10\ \text{TeV} \lesssim M_{DM} \lesssim 1\ \text{PeV}$
- $M_S \lesssim 10$ TeV
- $\delta_{lep} \sim 10^{-5}$ ⇒ leptogenesis is not fully resonant

Nicely predicted a signal at IceCube

(Anisimov,PDB,0812.5085;PDB, P.Ludl,S. Palomarez-Ruiz 1606.06238)

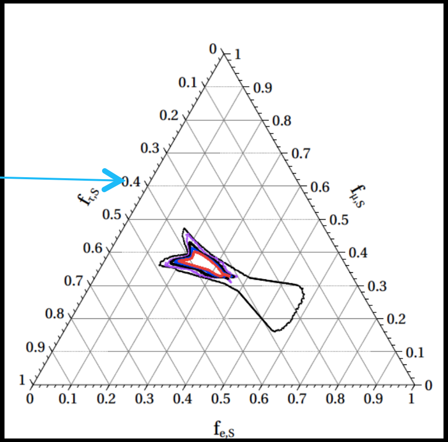
- DM neutrinos unavoidably decay today into A+leptons (A=H,Z,W) through the same mixing that produced them in the very early Universe
- Potentially testable high energy neutrino contribution

Energy neutrino flux



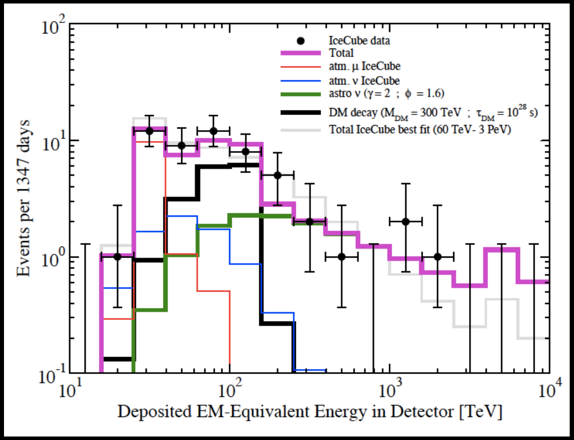
Hard component

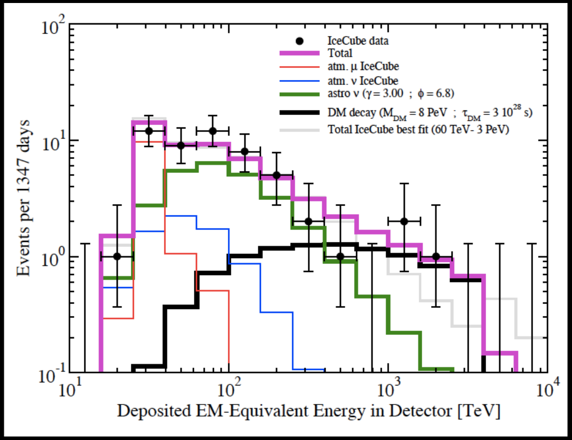
Flavour composition at the detector



Neutrino events at IceCube: 2 examples

M_{DM}=300TeV



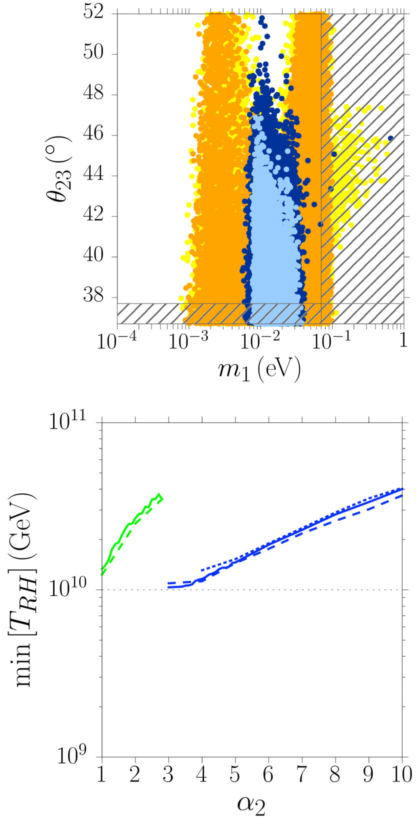


M_{DM}=8 PeV

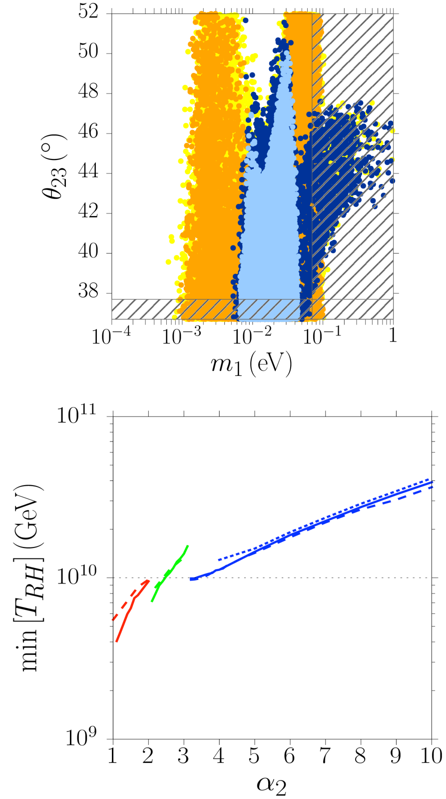
SUSY SO(10)-inspired leptogenesis

(PDB, Re Fiorentin, Marzola, 1512.06739)

tan β=5



tan β=50



It is possible to lower T_{RH} to values consistent with the gravitino problem for $m_g \gtrsim 30$ TeV

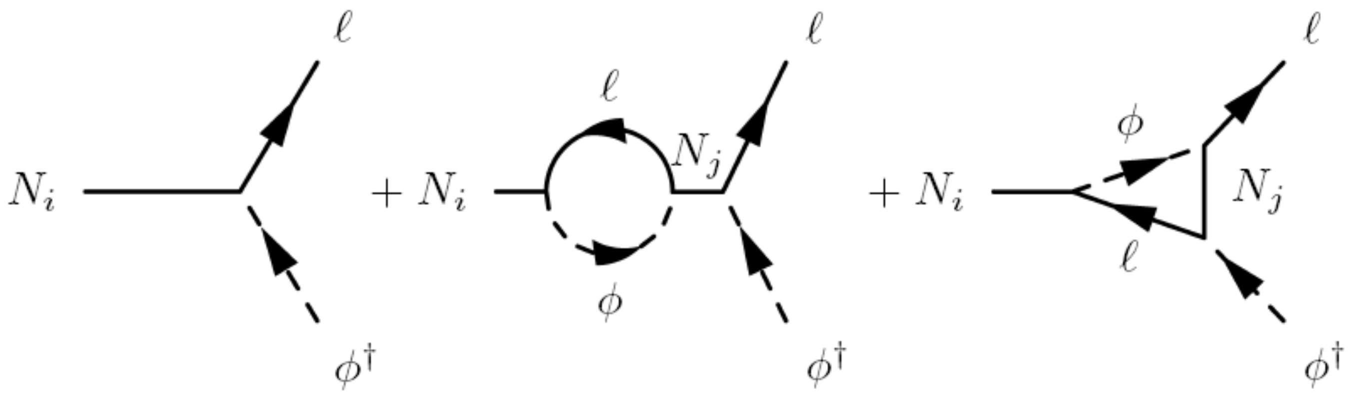
(Kawasaki, Kohri, Moroi, 0804.3745)

Alternatively, for lower gravitino masses, one has to consider non-thermal SO(10)-inspired leptogenesis

(Blanchet, Marfatia 1006.2857)

Total CP asymmetries

(Flanz,Paschos,Sarkar'95; Covi,Roulet,Vissani'96; Buchmüller,Plümacher'98)



$$\varepsilon_i \simeq \frac{1}{8\pi\, v^2\, (m_D^\dagger\, m_D)_{ii}} \sum_{j\neq i} \mathrm{Im}\left[(m_D^\dagger\, m_D)^2_{ij}\right] \times \left[f_V\left(\frac{M_j^2}{M_i^2}\right) + f_S\left(\frac{M_j^2}{M_i^2}\right)\right]$$

It does not depend on U !

N_1 dominated leptogenesis

$$z \equiv \frac{M_1}{T}$$

$$\frac{dN_{N_1}}{dz} = -D_1\,\left(N_{N_1} - N_{N_1}^{\rm eq}\right)$$

$$\frac{dN_{B-L}}{dz} = -\varepsilon_1\,\frac{dN_{N_1}}{dz} - W_{ID}\,N_{B-L}$$

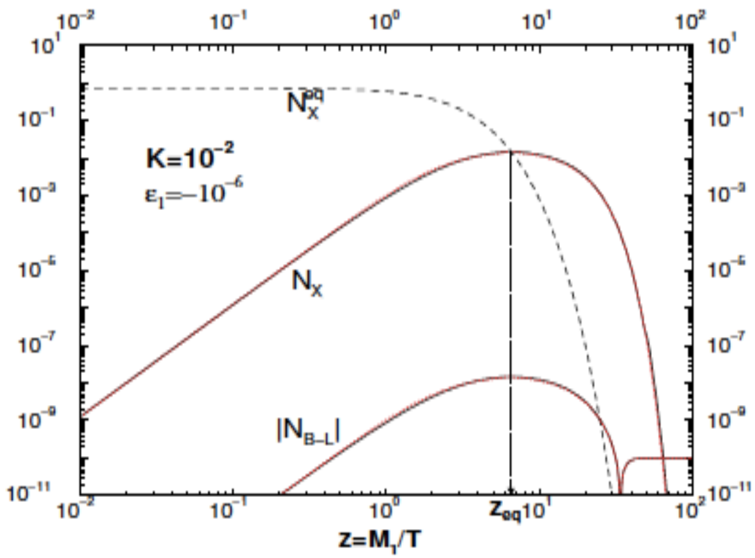
$$D_1 = \frac{\Gamma_{D,1}}{H\,z} = K_1\,z\,\left\langle\frac{1}{\gamma}\right\rangle\,, \quad W_{ID} \propto D_1 \propto K_1$$

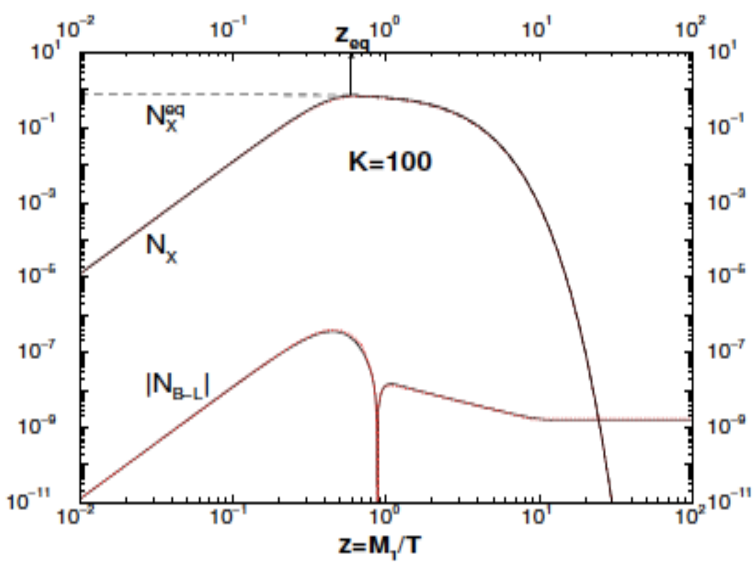
$$N_{B-L}(z;K_1,z_{\rm in}) = N_{B-L}^{\rm in}\,e^{-\int_{z_{\rm in}}^{z}\,dz'\,W_{ID}(z')} + \varepsilon_1\,\kappa_1(z)$$

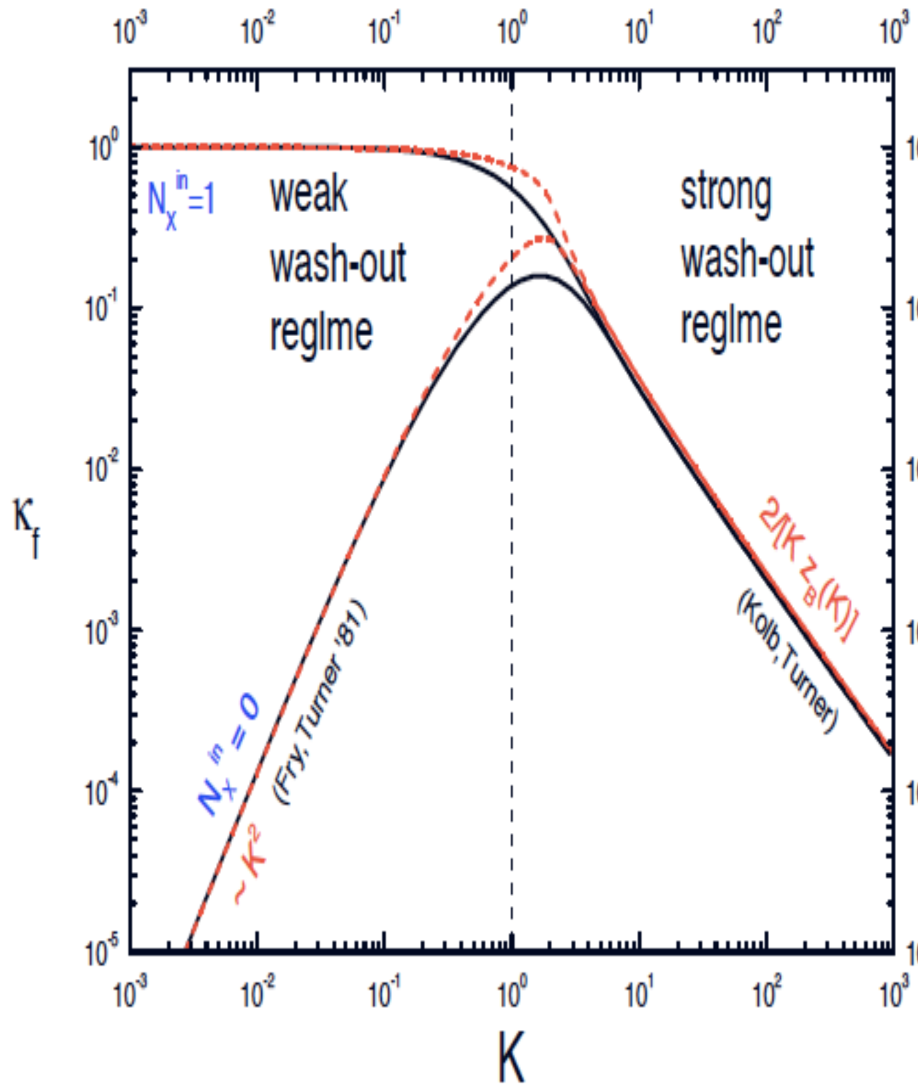
$$\kappa_1(z;K_1,z_{\rm in}) = -\int_{z_{\rm in}}^{z}\,dz'\,\left[\frac{dN_{N_1}}{dz'}\right]\,e^{-\int_{z'}^{z}\,dz''\,W_{ID}(z'')}$$

- Weak wash-out regime for $K_1 \lesssim 1$ (out-of-equilibrium picture recovered for $K_1 \rightarrow 0$)
- Strong wash-out regime for $K_1 \gtrsim 1$

Weak and strong wash-out: comparison





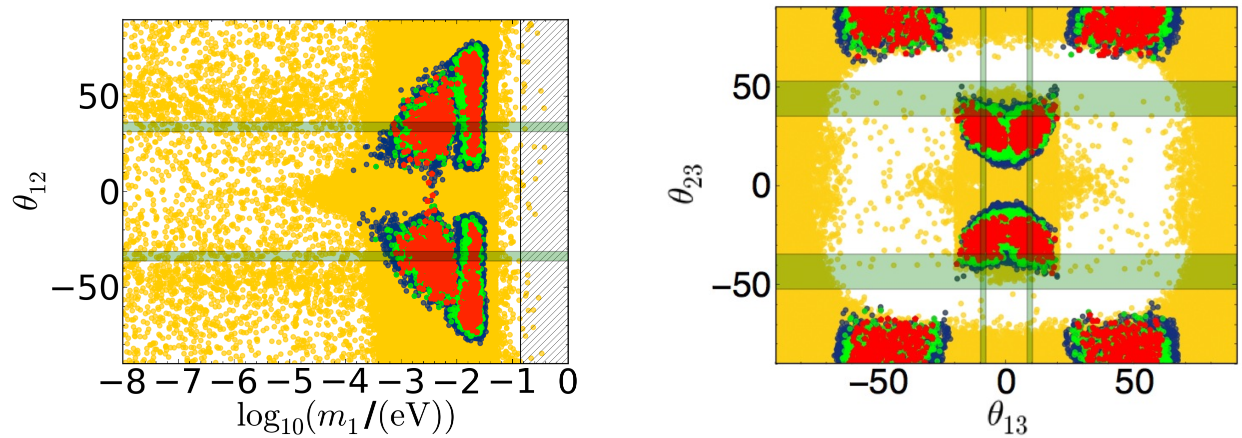


STSO10 solution: on the right track?

(PDB, Marzola '13)

What is the probability that the agreement is due to a coincidence?
This sets the statistical significance of the agreement

(N^p_{B-L}= 0, 0.001, 0.01, 0.1)



If the first octant is found then p≲10%

If NO is found then p≲5%

If sin δ < 0 is confirmed then p≲2%

If cos δ < 0 is found then p≲1% and then?

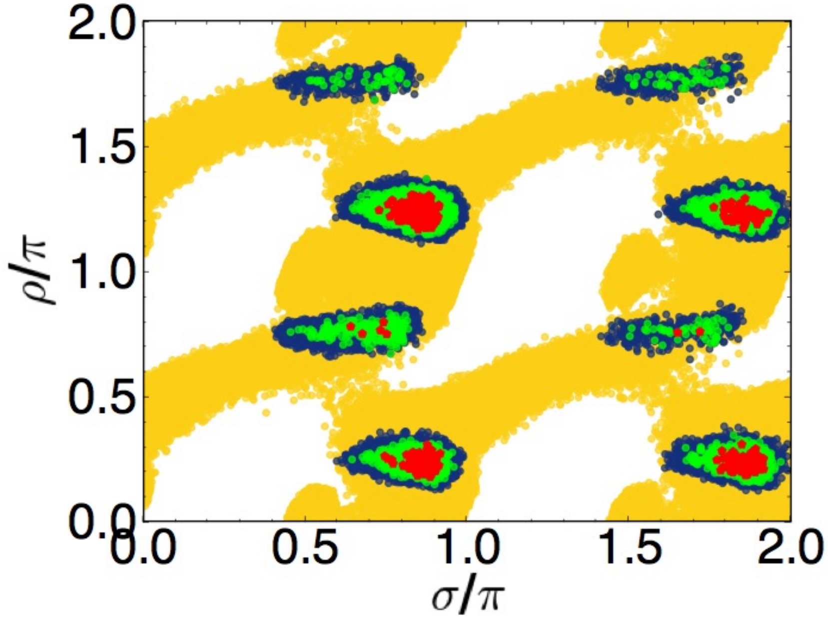
STSO10: Majorana phases and neutrinoless double beta decay

(PDB,Marzola1308.1107;PDB,Re Fiorentin,Marzola 1411.5478)

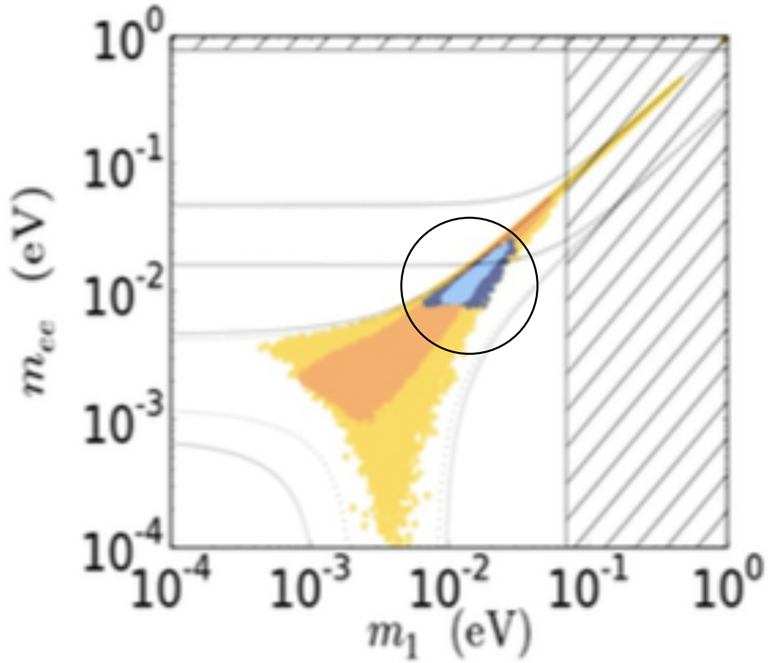
α_2=5

➢ NORMAL ORDERING

Majorana phases



$m_{ee} \simeq 0.8 m_1 \simeq 15$ meV



- Majorana phases are constrained around definite values
- Sharp prediction on the absolute neutrino mass scale: both on m_1 and m_{ee}
- Despite one has normal ordering, m_{ee} value might be within exp. Reach
- If also these predictions are satisfied exp, then $p \lesssim 0.01\%$ (conservative)

Decrypting SO(10)-inspired models

(Akhmedov, Frigerio, Smirnov, 2005; PDB, Re Fiorentin, Marzola, 1411.5478)

$$U_R \simeq \begin{pmatrix} 1 & -\frac{m_{D1}}{m_{D2}}\frac{m^\star_{\nu e\mu}}{m^\star_{\nu ee}} & \frac{m_{D1}}{m_{D3}}\frac{(m_\nu^{-1})^\star_{e\tau}}{(m_\nu^{-1})^\star_{\tau\tau}} \\ \frac{m_{D1}}{m_{D2}}\frac{m_{\nu e\mu}}{m_{\nu ee}} & 1 & \frac{m_{D2}}{m_{D3}}\frac{(m_\nu^{-1})^\star_{\mu\tau}}{(m_\nu^{-1})^\star_{\tau\tau}} \\ \frac{m_{D1}}{m_{D3}}\frac{m_{\nu e\tau}}{m_{\nu ee}} & -\frac{m_{D2}}{m_{D3}}\frac{(m_\nu^{-1})_{\mu\tau}}{(m_\nu^{-1})_{\tau\tau}} & 1 \end{pmatrix} D_\Phi \qquad D_\phi \equiv (e^{-i\frac{\Phi_1}{2}}, e^{-i\frac{\Phi_2}{2}}, e^{-i\frac{\Phi_3}{2}})$$

$$\boxed{M_1 \simeq \frac{m^2_{D1}}{|m_{\nu ee}|} \simeq \frac{\alpha_1^2\, m_u^2}{|m_{\nu ee}|} \simeq \alpha_1^2\, 10^5\,\mathrm{GeV}\left(\frac{m_u}{1\mathrm{MeV}}\right)^2\left(\frac{10\,\mathrm{meV}}{|m_{\nu ee}|}\right)}$$

$$\Phi_1 = \mathrm{Arg}[-m^\star_{\nu ee}]\,.$$

0νββ neutrino mass

$$\boxed{M_2 \simeq \frac{\alpha_2^2\, m_c^2}{m_1\, m_2\, m_3}\frac{|m_{\nu ee}|}{|(m_\nu^{-1})_{\tau\tau}|} \simeq \alpha_2^2\, 10^{11}\,\mathrm{GeV}\left(\frac{m_c}{400\mathrm{MeV}}\right)^2\left(\frac{|m_{\nu ee}|}{10\,\mathrm{meV}}\right)}$$

$$\Phi_2 = \mathrm{Arg}\left[\frac{m_{\nu ee}}{(m_\nu^{-1})_{\tau\tau}}\right] - 2\,(\rho+\sigma)$$

$$\boxed{M_3 \simeq \alpha_3^2\, m_t^2\, |(m_\nu^{-1})_{\tau\tau}| \simeq \alpha_3^2\, 10^{15}\,\mathrm{GeV}\left(\frac{m_t}{100\mathrm{GeV}}\right)^2\left(\frac{\mathrm{meV}}{m_1}\right)}\,.$$

$$\Phi_3 = \mathrm{Arg}[-(m_\nu^{-1})_{\tau\tau}]\,.$$

Decrypting SO(10)-inspired leptogenesis

(PDB, Re Fiorentin, Marzola, 1411.5478)

Finally, putting all together, one arrives to an expression for the final asymmetry:

$$\boxed{\begin{aligned} N_{B-L}^{\rm lep,f} &\simeq \frac{3}{16\,\pi}\,\frac{\alpha_2^2\,m_c^2}{v^2}\,\frac{|m_{\nu ee}|\,\left(|m_{\nu\tau\tau}^{-1}|^2+|m_{\nu\mu\tau}^{-1}|^2\right)^{-1}}{m_1\,m_2\,m_3}\,\frac{|m_{\nu\tau\tau}^{-1}|^2}{|m_{\nu\mu\tau}^{-1}|^2}\,\sin\alpha_L \\ &\times\ \kappa\left(\frac{m_1\,m_2\,m_3}{m_\star}\,\frac{|(m_\nu^{-1})_{\mu\tau}|^2}{|m_{\nu ee}|\,|(m_\nu^{-1})_{\tau\tau}|}\right) \\ &\times\ e^{-\frac{3\pi}{8}\,\frac{|m_{\nu e\tau}|^2}{m_\star\,|m_{\nu ee}|}}\,. \end{aligned}}$$

Effective SO(10)-inspired leptogenesis phase

$$\boxed{\alpha_L = {\rm Arg}\,[m_{\nu ee}] - 2\,{\rm Arg}[(m_\nu^{-1})_{\mu\tau}] + \pi - 2\,(\rho+\sigma)\,.}$$

<u>This analytical expression for the asymmetry fully reproduces all numerical constraints for $V_L=I$</u>

These results can be easily generalised to the case $V_L \neq 1$: all given expressions are still valid with the replacement: (Akhmedov, Frigerio, Smirnov, 2005)

$$\boxed{m_\nu \rightarrow \widetilde{m}_\nu \equiv \widetilde{V}_L\,m_\nu\,\widetilde{V}_L^T}$$

An example of realistic model:
SO(10)-inspired leptogenesis in the "A2Z model"

(S.F. King 2014)

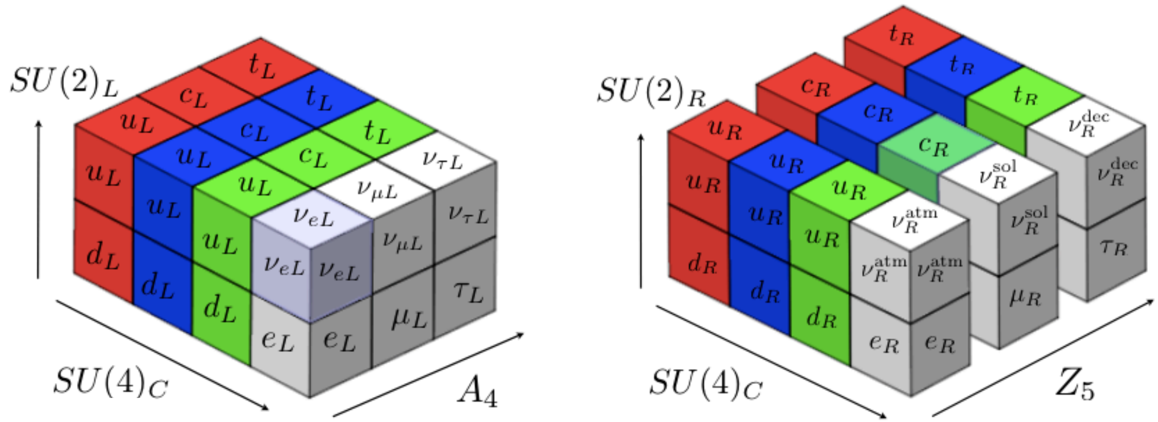


Figure 1: A to Z of flavour with Pati-Salam, where $A \equiv A_4$ and $Z \equiv Z_5$. The left-handed families form a triplet of A_4 and are doublets of $SU(2)_L$. The right-handed families are distinguished by Z_5 and are doublets of $SU(2)_R$. The $SU(4)_C$ unifies the quarks and leptons with leptons as the fourth colour, depicted here as white.

Neutrino sector:

$$Y'^{\nu}_{LR} = \begin{pmatrix} 0 & be^{-i3\pi/5} & 0 \\ ae^{-i3\pi/5} & 4be^{-i3\pi/5} & 0 \\ ae^{-i3\pi/5} & 2be^{-i3\pi/5} & ce^{i\phi} \end{pmatrix}, \quad M'_R = \begin{pmatrix} M'_{11}e^{2i\xi} & 0 & M'_{13}e^{i\xi} \\ 0 & M'_{22}e^{i\xi} & 0 \\ M'_{13}e^{i\xi} & 0 & M'_{33} \end{pmatrix}$$

CASE A:

$$m^D_{\nu 1} = m_{\rm up}, \quad m^D_{\nu 2} = m_{\rm charm}, \quad m^D_{\nu 3} = m_{\rm top}$$

CASE B:

$$m^D_{\nu 1} \approx m_{\rm up}, \quad m^D_{\nu 2} \approx 3\, m_{\rm charm}, \quad m^D_{\nu 3} \approx \frac{1}{3}\, m_{\rm top}$$

There are 2 solutions (only for NO)

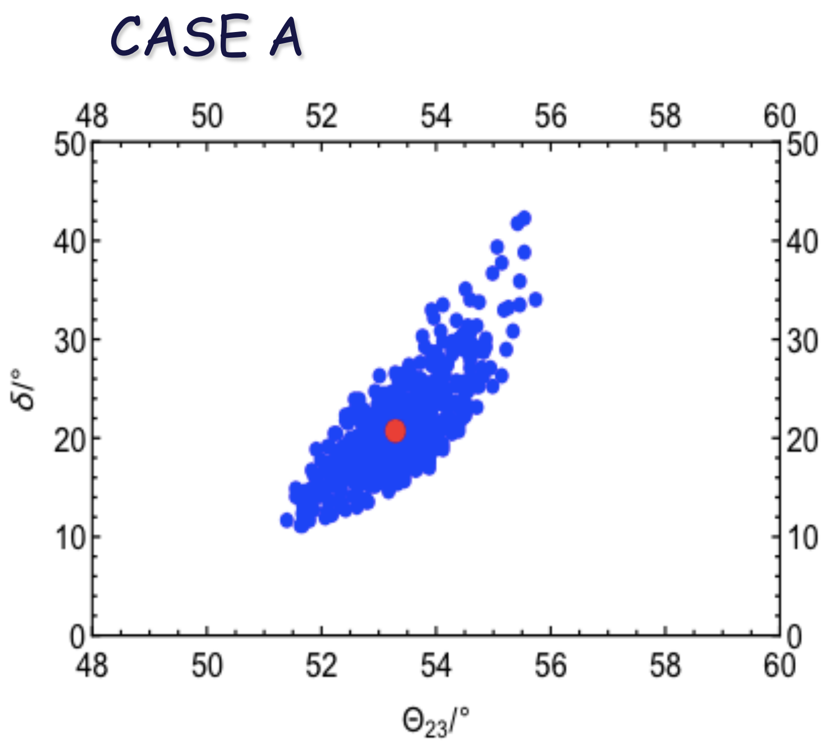
(PDB, S.King 1507.06431)

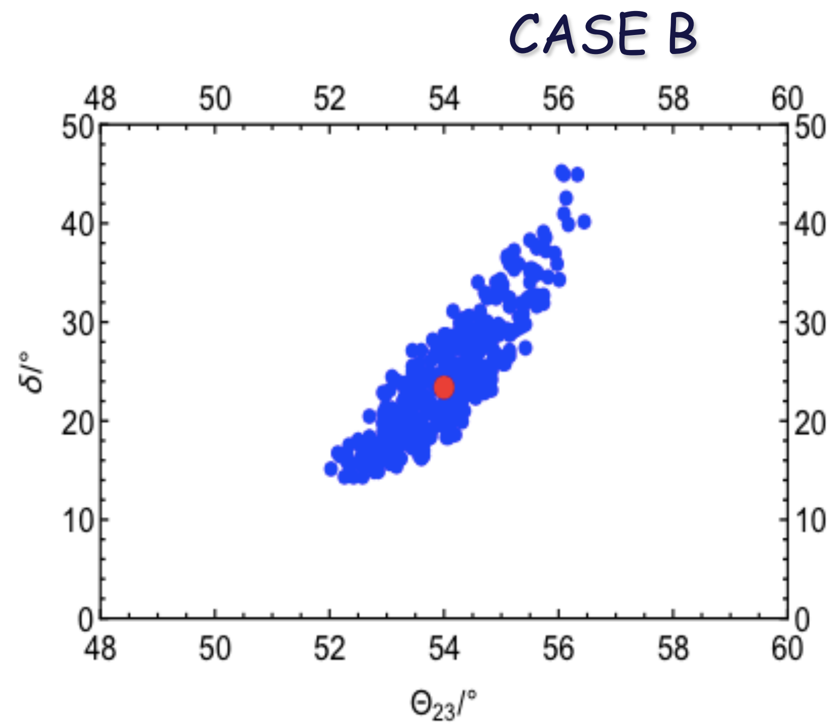
The spectrum
is not so strongly
hierarchical:
it is in the proximity
of crossing level
solutions

CASE	A	B
ξ	$+4\pi/5$	
χ^2_{min}	5.15	6.1
$M_1/10^7\mathrm{GeV}$	15	2.7
$M_2/10^{10}\mathrm{GeV}$	0.483	4.35
$M_3/10^{12}$ GeV	2.16	1.31
$\lvert\gamma\rvert$	203	38
m_1/meV	2.3	2.3
m_2/meV ($p_{\Delta m^2_{12}}$)	8.93 (-0.22)	8.94 (-0.25)
m_3/meV ($p_{\Delta m^2_{13}}$)	49.7 (+0.17)	49.7 (+0.21)
$\sum_i m_i/\mathrm{meV}$	61	61
m_{ee}/meV	1.95	1.95
$\theta_{12}/^\circ$ ($p_{\theta_{12}}$)	33.0 (-0.58)	33.0 (-0.66)
$\theta_{13}/^\circ$ ($p_{\theta_{13}}$)	8.40 (-0.47)	8.40 (-0.49)
$\theta_{23}/^\circ$ ($p_{\theta_{23}}$)	53.3 (+2.1)	54.0 (+2.3)
$\delta/^\circ$	20.8	23.5
$\eta_B/10^{-10}$ (p_{η_B})	6.101 (+0.01)	6.101 (+0.01)
$\varepsilon_{2\tau}$	-8.1×10^{-6}	-1.3×10^{-5}
$K_{1\mu}$	0.11	0.58
$K_{1\tau}$	4341	800
$K_{2\tau}$	7.3	7.3
$K_{2\mu}$	29.2	29.2
K_{2e}	1.8	1.8

There are 2 solutions (only for NO)

(PDB, S.F. King 1507.06431)





This region will be tested relatively quickly

Quantifying the fine-tuning

(PDB, S.King 2015)

Analytical expression also for the orthogonal matrix:

$$\Omega \simeq \begin{pmatrix} -\frac{\sqrt{m_1\,|\widetilde{m}_{\nu 11}|}}{\widetilde{m}_{\nu 11}}\,U_{e1} & \sqrt{\frac{m_2\,m_3\,|(\widetilde{m}_\nu^{-1})_{33}|}{|\widetilde{m}_{\nu 11}|}}\left(U^\star_{\mu 1} - U^\star_{\tau 1}\,\frac{(\widetilde{m}_\nu^{-1})_{23}}{(\widetilde{m}_\nu^{-1})_{33}}\right) & \frac{U^\star_{31}}{\sqrt{m_1\,|(\widetilde{m}_\nu^{-1})_{33}|}} \\ -\frac{\sqrt{m_2\,|\widetilde{m}_{\nu 11}|}}{\widetilde{m}_{\nu 11}}\,U_{e2} & \sqrt{\frac{m_1\,m_3\,|(\widetilde{m}_\nu^{-1})_{33}|}{|\widetilde{m}_{\nu 11}|}}\left(U^\star_{\mu 2} - U^\star_{\tau 2}\,\frac{(\widetilde{m}_\nu^{-1})_{23}}{(\widetilde{m}_\nu^{-1})_{33}}\right) & \frac{U^\star_{32}}{\sqrt{m_2\,|(\widetilde{m}_\nu^{-1})_{33}|}} \\ -\frac{\sqrt{m_3\,|\widetilde{m}_{\nu 11}|}}{\widetilde{m}_{\nu 11}}\,U_{e3} & \sqrt{\frac{m_1\,m_2\,|(\widetilde{m}_\nu^{-1})_{33}|}{|\widetilde{m}_{\nu 11}|}}\left(U^\star_{\mu 3} - U^\star_{\tau 3}\,\frac{(\widetilde{m}_\nu^{-1})_{23}}{(\widetilde{m}_\nu^{-1})_{33}}\right) & \frac{U^\star_{33}}{\sqrt{m_3\,|(\widetilde{m}_\nu^{-1})_{33}|}} \end{pmatrix} D_\Phi\,,$$

$$\Omega^{(\mathrm{CASEA})} \simeq \begin{pmatrix} -4.40016 - 15.9889\,i & 0.0930875 - 0.894045\,i & -16.0396 + 4.38107\,i \\ -15.9446 + 3.40333\,i & -1.15394 + 0.0537137\,i & 3.40494 + 15.9553\,i \\ -3.69174 + 4.35811\,i & 0.709793 + 0.204576\,i & 4.37787 + 3.64191\,i \end{pmatrix}$$

$$\Omega^{(CASEB)} \simeq \begin{pmatrix} -1.77835 - 6.85986\,i & 0.108413 - 0.897431\,i & -6.97828 + 1.73423\,i \\ -6.87598 + 1.34103\,i & -1.15331 + 0.0386159\,i & 1.34278 + 6.90018\,i \\ -1.64314 + 1.81259\,i & 0.710523 + 0.199612\,i & 1.85785 + 1.52677\,i \end{pmatrix}$$

- Fine tuned cancellations in the see-saw formula at the level of $|\Omega_{ij}|^{-2}$ this seems to be quite a recurrent issue in fits…..

A popular class of SO(10) models

(Fritzsch, Minkowski, Annals Phys. 93 (1975) 193-266; R.Slansky, Phys.Rept. 79 (1981) 1-128; G.G. Ross, GUTs, 1985; Dutta, Mimura, Mohapatra, hep-ph/0507319; G. Senjanovic hep-ph/0612312)

In SO(10) models each SM particles generation + 1 RH neutrino are assigned to a single 16-dim representation. Masses of fermions arise from Yukawa interactions of two 16s with vevs of suitable Higgs fields. Since:

$$16 \otimes 16 = 10_S \oplus \overline{126}_S \oplus 120_A;$$

The Higgs fields of renormalizable SO(10) models can belong to 10-, 126-,120-dim representations yielding Yukawa part of the Lagrangian

$$\mathcal{L}_Y = 16\,(Y_{10} 10_H + Y_{126}\overline{126}_H + Y_{120} 120_H)\,16\,.$$

After SSB of the fermions at M_{GUT}=2x10^{16} GeV one obtains the masses:

up-quark mass matrix	$M_u = \boxed{v_{10}^u Y_{10}} + v_{126}^u Y_{126} + v_{120}^u Y_{120}\,,$
down-quark mass matrix	$M_d = \boxed{v_{10}^d Y_{10}} + v_{126}^d Y_{126} + v_{120}^d Y_{120}\,,$
neutrino mass matrix	$M_D = \boxed{v_{10}^u Y_{10}} - 3v_{126}^u Y_{126} + v_{120}^D Y_{120}\,,$
charged lepton mass matrix	$M_l = \boxed{v_{10}^d Y_{10}} - 3v_{126}^d Y_{126} + v_{120}^l Y_{120}\,,$
RH neutrino mass matrix	$M_R = v_{126}^R Y_{126}\,,$
LH neutrino mass matrix	$M_L = v_{126}^L Y_{126}\,,$

Simplest case but clearly non-realistic: it predicts no mixing at all (both in quark and lepton Sectors). For realistic models one has to add at least the 126 contribution

NOTE: these models do respect SO(10)-inspired conditions

Recent fits within SO(10) models

(Joshipura Patel 2011; Rodejohann, Dueck '13)

No type II seesaw contribution: it does not seem to help the fits

Minimal Model with $10_H + \overline{126}_H$ (MN, MS)

"full" Higgs Content $10_H + \overline{126}_H + 120_H$ (FN, FS)

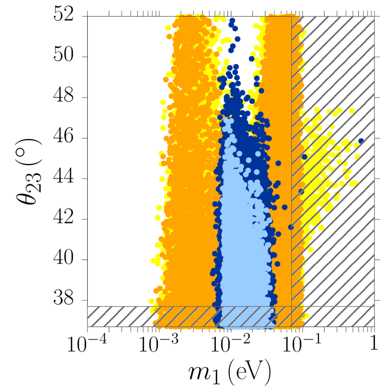
Mod	Comments	$\langle m_\nu \rangle$ [meV]	δ^l_{CP} [rad]	$\sin^2 \theta^l_{23}$	m_0 [meV]	M_3 [GeV]	M_2 [GeV]	M_1 [GeV]	$\chi^2_{\min}$
MN	no RGE, NH	0.35	0.7	0.406	3.03	5.5×10^{12}	7.2×10^{11}	1.5×10^{10}	1.10
MN	RGE, NH	0.49	6.0	0.346	2.40	3.6×10^{12}	2.0×10^{11}	1.2×10^{11}	23.0
MS	no RGE, NH	0.38	0.27	0.387	2.58	3.9×10^{12}	7.2×10^{11}	1.6×10^{10}	9.41
MS	RGE, NH	0.44	2.8	0.410	6.83	1.1×10^{12}	5.7×10^{10}	1.5×10^{10}	3.29
FN	no RGE, NH	4.96	1.7	0.410	8.8	1.9×10^{13}	2.8×10^{12}	2.2×10^{10}	6.6×10^{-5}
FN	RGE, NH	2.87	5.0	0.410	1.54	9.9×10^{14}	7.3×10^{13}	1.2×10^{13}	11.2
FS	no RGE, NH	0.75	0.5	0.410	1.16	1.5×10^{13}	5.3×10^{11}	5.7×10^{10}	9.0×10^{-10}
FS	RGE, NH	0.78	5.4	0.410	3.17	4.2×10^{13}	4.9×10^{11}	4.9×10^{11}	6.9×10^{-6}
FN	no RGE, IH	35.37	5.4	0.590	35.85	2.2×10^{13}	4.9×10^{12}	9.2×10^{11}	2.5×10^{-4}
FN	RGE, IH	35.52	4.7	0.590	30.24	1.1×10^{13}	3.5×10^{12}	5.5×10^{11}	13.3
FS	no RGE, IH	44.21	0.3	0.590	6.27	1.2×10^{13}	4.2×10^{11}	3.5×10^{7}	3.9×10^{-8}
FS	RGE, IH	24.22	3.6	0.590	11.97	1.2×10^{13}	3.1×10^{11}	2.0×10^{3}	0.602

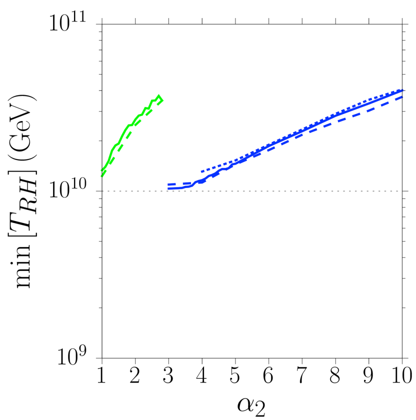
Recently Fong,Meloni,Meroni,Nardi(1412.4776) have included leptogenesis for the non-SUSY case obtaining successful leptogenesis: but such a compact RN neutrino spectrum implies huge fine-tuning. Too simplistic models? What solution: non renormalizable terms? Type II seesaw term? SUSY seems to improve the fits and also give 1 hier. solution

SUSY SO(10)-inspired leptogenesis

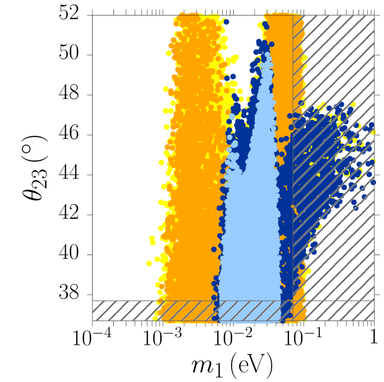
(PDB, Re Fiorentin, Marzola, 1512.06739)

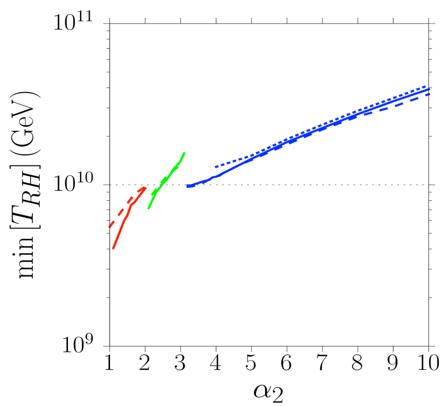
tan β=5





tan β=50





It is possible to lower T_{RH} to values consistent with the gravitino problem for $m_g \gtrsim 30$ TeV

(Kawasaki, Kohri, Moroi, 0804.3745)

Alternatively, for lower gravitino masses, one has to consider non-thermal SO(10)-inspired leptogenesis

(Blanchet, Marfatia 1006.2857)

Leptogenesis in the "A2Z model"

(PDB, S.King 2015)

The only sizeable CP asymmetry is the tauon asymmetry but $K_{1t} >> 1$!

Flavour coupling (mainly due to the hypercharge Higgs asymmetry) is then crucial to produce the correct asymmetry:
(Antusch,PDB,Jones,King 2011)

$$\eta_B \simeq \sum_{\alpha=e,\mu,\tau} \eta_B^{(\alpha)} \,, \qquad \eta_B^{(\tau)} \simeq 0.01\, \varepsilon_{2\tau}\, \kappa(K_{2\tau})\, e^{-\frac{3\pi}{8}\, K_{1\tau}}$$

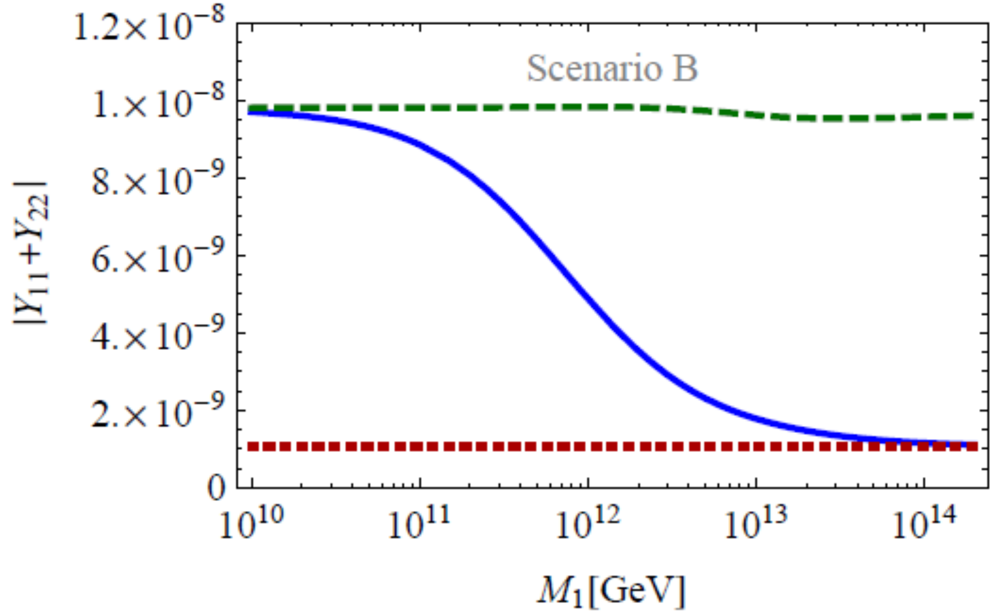
$$\eta_B^{(e)} \simeq -0.01\, \varepsilon_{2\tau}\, \kappa(K_{2\tau})\, \frac{K_{2e}}{K_{2e}+K_{2\mu}}\, C^{(2)}_{\tau^{\perp}\tau}\, e^{-\frac{3\pi}{8}\, K_{1e}}$$

$$\eta_B^{(\mu)} \simeq -\left(\frac{K_{2\mu}}{K_{2e}+K_{2\mu}}\, C^{(2)}_{\tau^{\perp}\tau} - \frac{K_{1\mu}}{K_{1\tau}}\, C^{(3)}_{\mu\tau}\right)\, e^{-\frac{3\pi}{8}\, K_{1\mu}} \,.$$

Density matrix and CTP formalism to describe the transition regimes

(De Simone, Riotto '06; Beneke, Gabrecht, Fidler, Herranen, Schwaller '10)

$$\frac{dY_{\alpha\beta}}{dz} = \frac{1}{sz\,H(z)}\left[(\gamma_D + \gamma_{\Delta L=1})\left(\frac{Y_{N_1}}{Y_{N_1}^{eq}} - 1\right)\epsilon_{\alpha\beta} - \frac{1}{2Y_\ell^{eq}}\left\{\gamma_D + \gamma_{\Delta L=1}, Y\right\}_{\alpha\beta}\right] - \left[\sigma_2 \mathrm{Re}(\Lambda) + \sigma_1 |\mathrm{Im}(\Lambda)|\right] Y_{\alpha\beta}$$

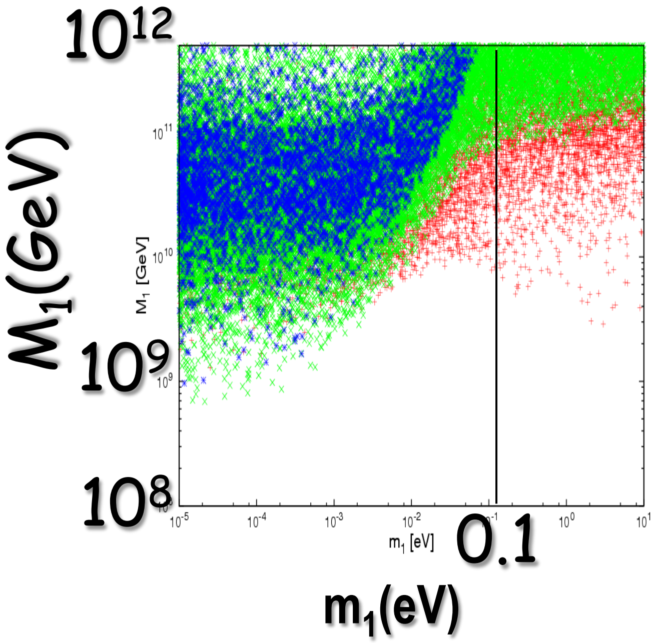


Fully two-flavoured regime limit

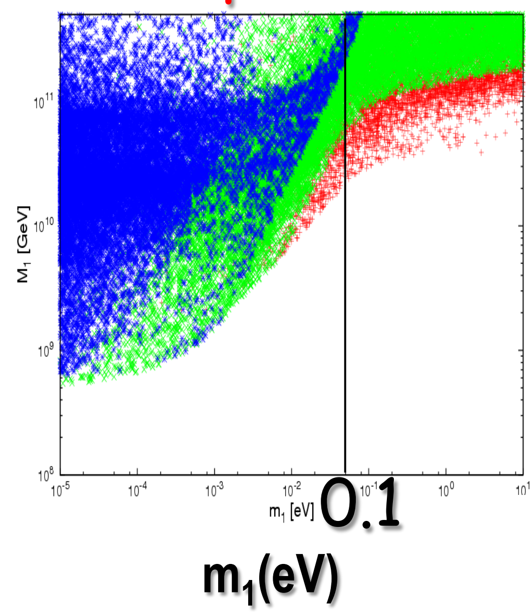
Unflavoured regime limit

Neutrino mass bounds and role of PMNS phases

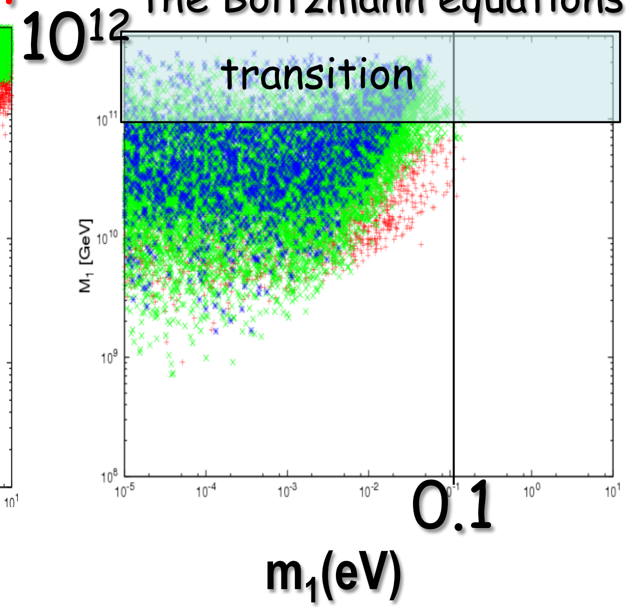
(Abada et al.' 07; Blanchet,PDB,Raffelt;Blanchet,PDB '08)

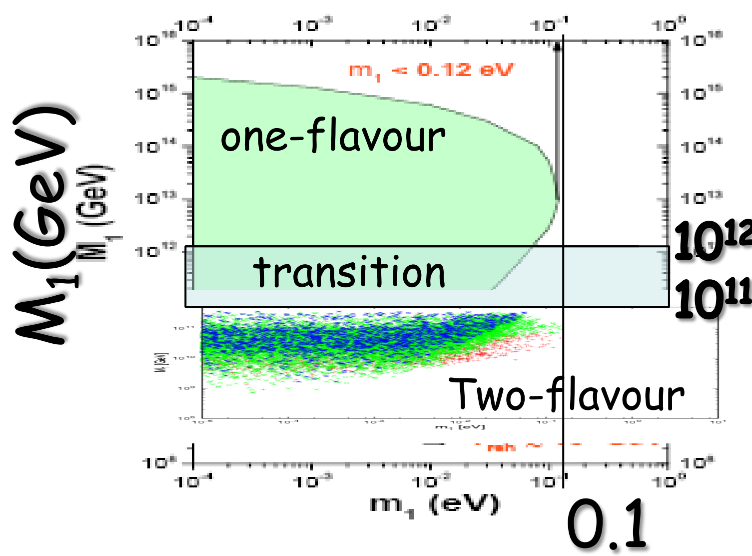


PMNS phases off



Imposing the validity of the Boltzmann equations



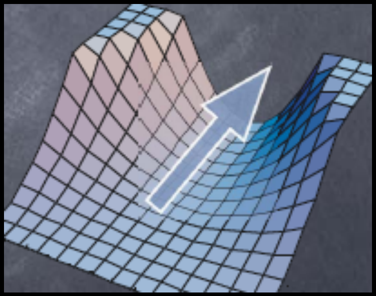


Affleck-Dine Baryogenesis

(Affleck, Dine '85)

In the Supersymmetric SM there are many "flat directions" in the space of a field composed of squarks and/or sleptons

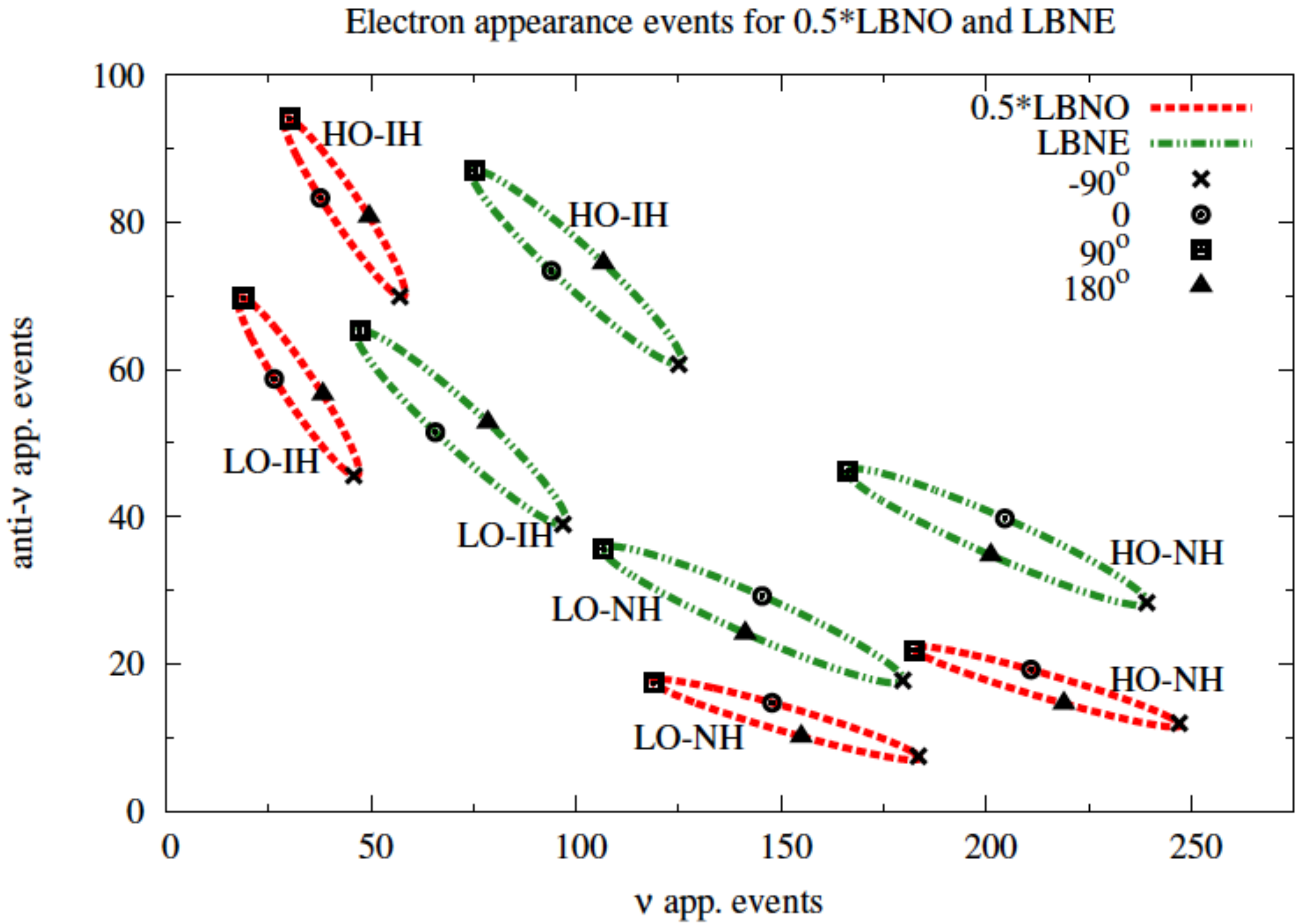
$$V(\phi) = \underbrace{\sum_i \left| \frac{\partial W}{\partial \phi_i} \right|^2}_{\text{F term}} + \underbrace{\frac{1}{2} \sum_A \left(\sum_{ij} \phi_i^* (t_A)_{ij} \phi_j \right)^2}_{\text{D term}}$$



A flat direction can be parametrized in terms of a complex field (AD field) that carries a baryon number that is violated dynamically during inflation

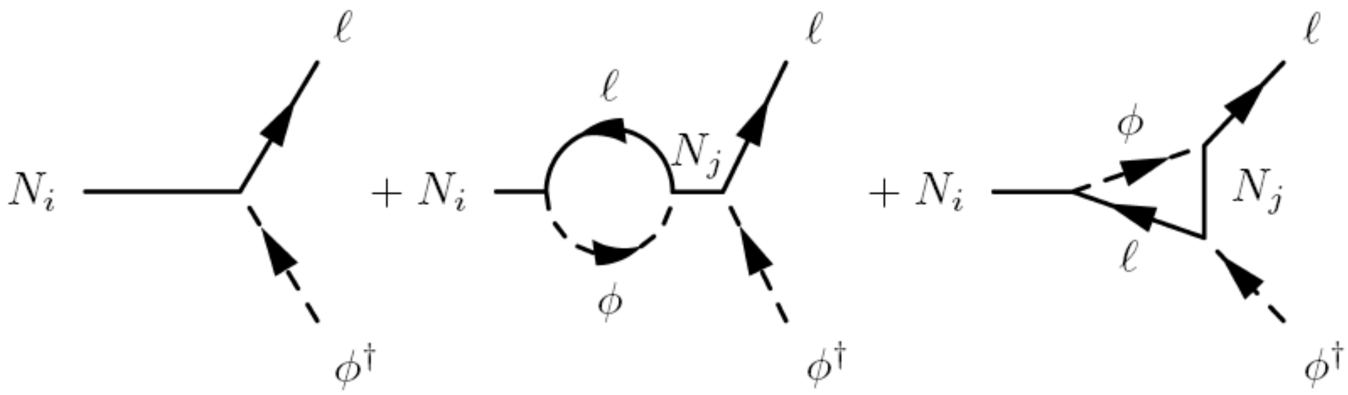
$$\frac{n_B}{s} \sim 10^{-10} \left(\frac{m_{3/2}}{m_\Phi} \right) \left(\frac{m_\Phi}{\mathrm{TeV}} \right)^{-\frac{1}{2}} \left(\frac{M}{M_P} \right)^{\frac{3}{2}} \left(\frac{T_R}{10\,\mathrm{GeV}} \right)$$

The final asymmetry is ? T_{RH} and the observed one can be reproduced for low values T_{RH} ? 10 GeV !


Electron appearance events for 0.5*LBNO and LBNE
0.5*LBNO
LBNE
-90°
0
90°
180°
HO-IH
HO-IH
LO-IH
LO-IH
LO-NH
LO-NH
HO-NH
HO-NH
anti-ν app. events
ν app. events
0
20
40
60
80
100
0
50
100
150
200
250

Total CP asymmetries

(Flanz,Paschos,Sarkar'95; Covi,Roulet,Vissani'96; Buchmüller,Plümacher'98)



$$\varepsilon_i \simeq \frac{1}{8\pi v^2 (m_D^\dagger m_D)_{ii}} \sum_{j\neq i} \mathrm{Im}\left[(m_D^\dagger m_D)^2_{ij}\right] \times \left[f_V\left(\frac{M_j^2}{M_i^2}\right) + f_S\left(\frac{M_j^2}{M_i^2}\right)\right]$$

It does not depend on U !

Additional contribution to CP violation:

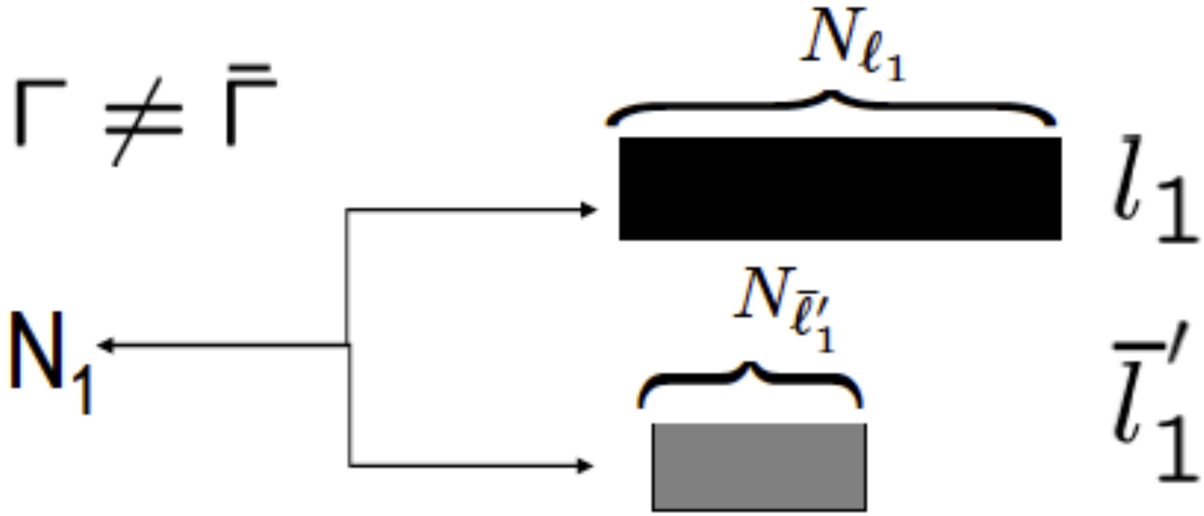
(Nardi, Racker, Roulet '06)

($\alpha = \tau,\ e{+}\mu$)

$$\boxed{\varepsilon_{1\alpha} = P^0_{1\alpha}\,\varepsilon_1 + \frac{\Delta P_{1\alpha}}{2}}$$

depends on U !

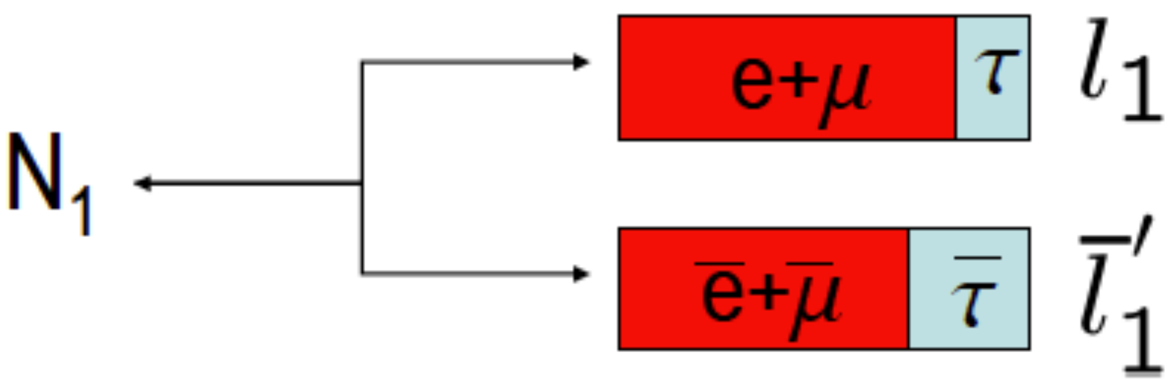
1) $\Gamma \neq \bar{\Gamma}$



$\Rightarrow\ P^0_{1\alpha}\,\varepsilon_1$

2) $|\bar{l}'_1\rangle \neq CP|l_1\rangle$

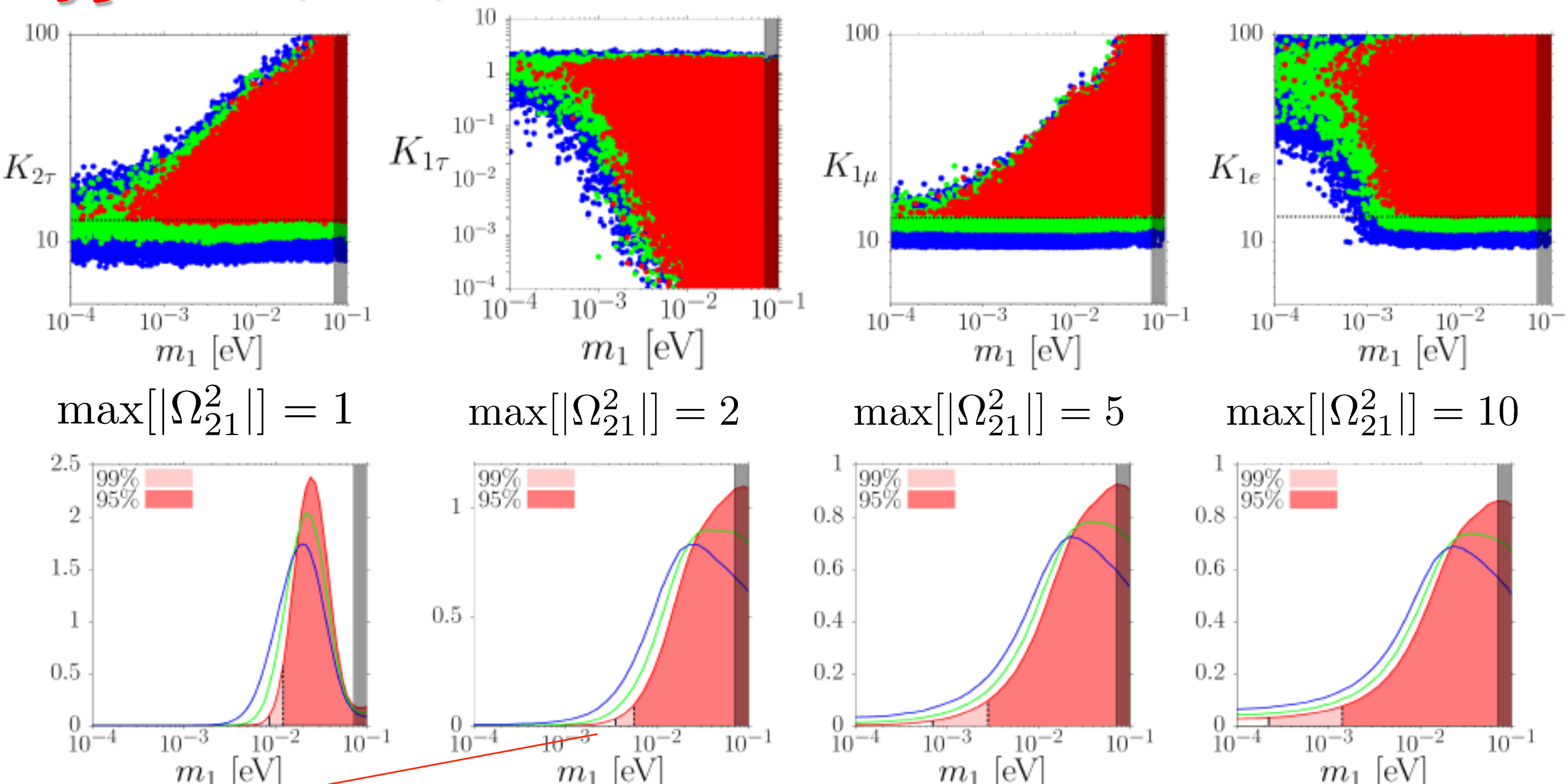
+



$\Rightarrow\ \dfrac{\Delta P_{1\alpha}}{2}$

A lower bound on neutrino masses (IO)

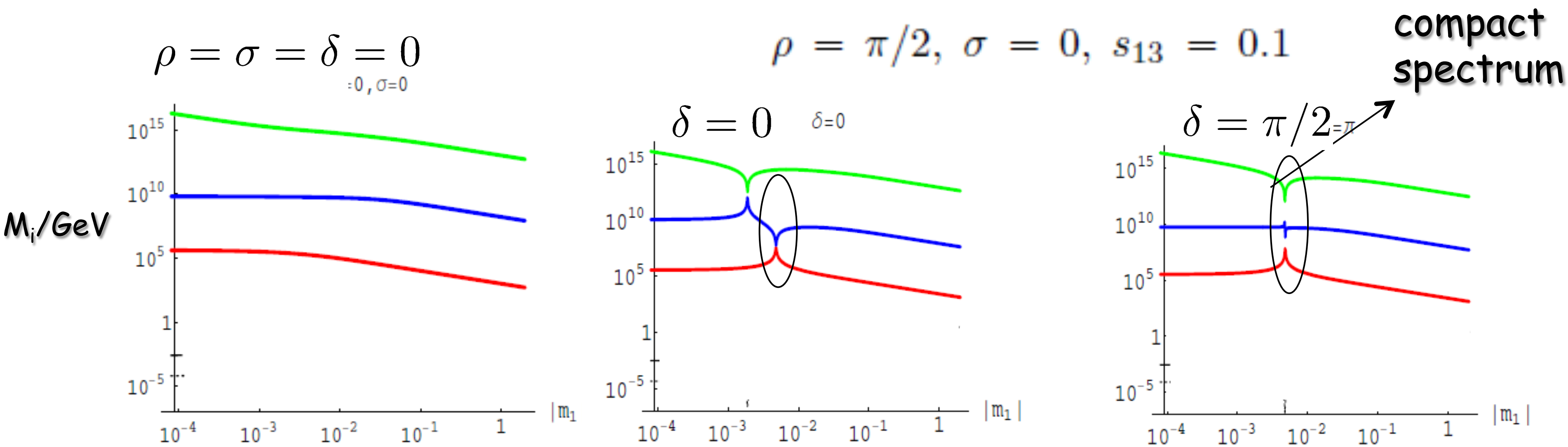
$N^{P,i}_{B-L}$ = 0.001, 0.01, 0.1 $\max[|\Omega^2_{21}|] = 2$ INVERTED ORDERING



$m_1 \gtrsim 3$ meV $\Rightarrow \Sigma_i m_i \gtrsim 100$ meV (not necessarily deviation from HL)

Crossing level solutions

(Akhmedov, Frigerio, Smirnov hep-ph/0305322)



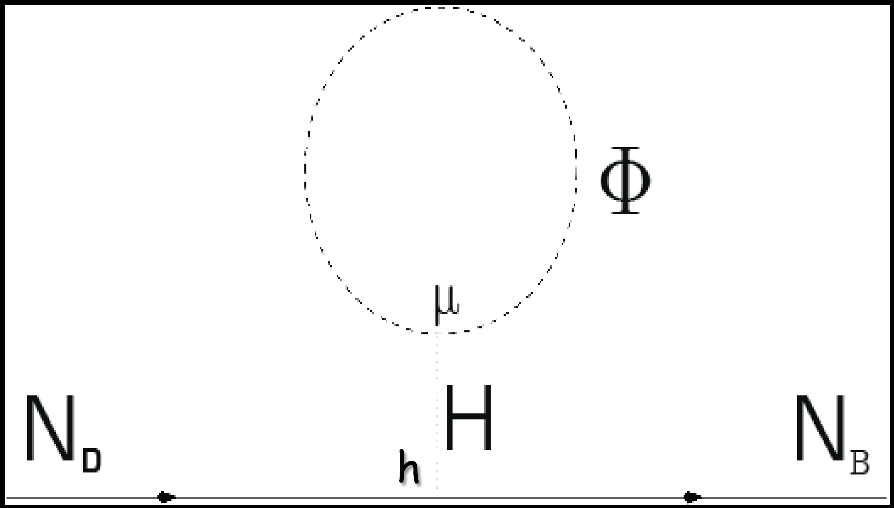
- About the crossing levels the N_1 CP asymmetry is enhanced

- The correct BAU can be attained for a fine tuned choice of parameters: many realistic models have made use of these solutions

(e.g. Ji, Mohapatra,Nasri '10; Buccella, Falcone, Nardi, '12; Altarelli, Meloni '14, Feng, Meloni, Meroni, Nardi '15; Addazi, Bianchi, Ricciardi 1510.00243)

A possible GUT origin

(Anisimov,PDB, 2010, unpublished)

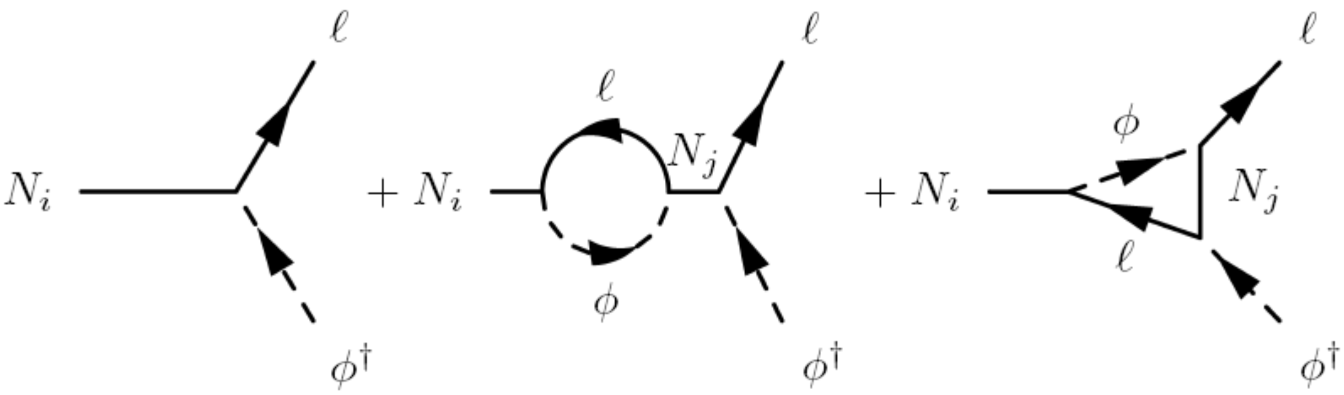


$$\frac{1}{\Lambda_{\rm eff}} = \frac{h\mu}{M_{\rm GUT}^2}$$

$$\Lambda_{\rm eff} >> M_{\rm GUT}\,!$$

Total CP asymmetries

(Flanz, Paschos, Sarkar'95; Covi, Roulet, Vissani'96; Buchmüller, Plümacher'98)



$$\varepsilon_i \simeq \frac{1}{8\pi v^2 (m_D^\dagger m_D)_{ii}} \sum_{j\neq i} \mathrm{Im}\left[(m_D^\dagger m_D)^2_{ij}\right] \times \left[f_V\left(\frac{M_j^2}{M_i^2}\right) + f_S\left(\frac{M_j^2}{M_i^2}\right)\right]$$

It does not depend on U !

2 fully flavoured regime

Flavoured decay parameters:

$$K_{1\alpha} = P^0_{1\alpha} K_1 = \left|\sum_k \sqrt{\frac{m_k}{m_*}} U_{\alpha k} \Omega_{k1}\right|^2 \leq K_1 \qquad \sum_\alpha K_{1\alpha} = K_1$$

$$\Rightarrow \boxed{\varepsilon_{1\alpha} \equiv -\frac{P_{1\alpha}\Gamma_1 - \bar{P}_{1\alpha}\bar{\Gamma}_1}{\Gamma_1+\bar{\Gamma}_1} = P^0_{1\alpha}\,\varepsilon_1 + \Delta P_{1\alpha}(\Omega, U)/2}$$

($\alpha = \tau$, e+μ)

$$\Rightarrow \boxed{N^{\rm fin}_{B-L} = \sum_\alpha \varepsilon_{1\alpha}\,\kappa^{\rm fin}_{1\alpha} \simeq 2\,\varepsilon_1\,\kappa_1^{\rm fin} + \frac{\Delta P_{1\alpha}}{2}\left[\kappa^{\rm f}(K_{1\alpha}) - \kappa^{\rm fin}(K_{1\beta})\right]}$$

Flavour effects introduce an explicit dependence on U and can in some case greatly enhance the asymmetry compared to the unflavoured case.

3 MAIN APPLICATIONS AND CONSEQUENCES OF FLAVOUR EFFECTS:

- Lower bound on M_1 (an therefore on T_{RH}) is not relaxed
  upper bound on m_1 is slightly relaxed to ~0.2eV but if wash-out is strong then Low energy phases can strongly affect the final asymmetry (second term)

- In the case of real Ω ⇒ all CP violation stems from low energy phases; if also Majorana phases are CP conserving only δ would be responsible for the asymmetry: ⇒ DIRAC PHASE LEPTOGENESIS: $\eta_{B0} \propto |\sin\delta|\,\sin\Theta_{13}$

- Asymmetry produced from heavier RH neutrinos also contributes to the asymmetry and has to be taken into account:
  IT OPENS NEW INTERESTING OPPORTUNITIES

Remarks on the role of δ in leptogenesis

Dirac phase leptogenesis:

- It could work but only for $M_1 \gtrsim 5\times10^{11}$ GeV (plus other conditions on Ω)
  $\Rightarrow$density matrix calculation needed!

- No reasons for Ω to be real except when it is a permutation of identity (from discrete flavour models) but then all CP asymmetries vanish! So one needs quite a a special Ω

- In general the contribution from δ is *overwhelmed* by the high energy phases in Ω

General considerations:

- CP violating value of δ is strictly speaking neither necessary nor sufficient condition for successful leptogenesis and no specific value is favoured model independently but….

- ….it is important to exclude CP conserving values since from $m_D = U\sqrt{D_m}\,\Omega\sqrt{D_M}$ one expects for generic m_D that if there are phases in U then there are also phases in Ω, vice-versa if there are no phases in U one might suspect that also Ω is real (disaster!):

  discovering CP violating value of δ would support a complex m_D

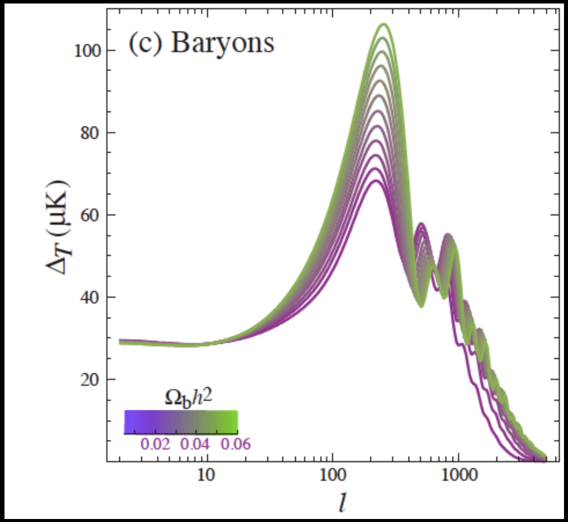
Seesaw parameter space

A parameter reduction would help and can occur in various ways:

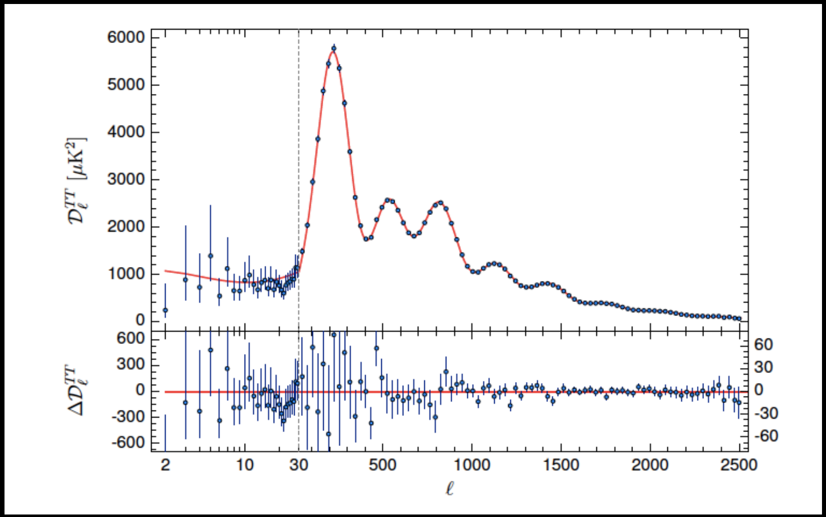
- $n_B = n_B^{CMB}$ is satisfied around "peaks";
- some parameters cancel out in the asymmetry calculation;
- imposing independence of the initial conditions (strong thermal leptog.);
- imposing model dependent conditions on m_D (e.g. SO(10)-inspired)
- additional phenomenological constraints (e.g. Dark Matter)

Baryon asymmetry of the universe

(Hu, Dodelson, astro-ph/0110414)



(Planck 2018, 1807.06209)



(68% CL,TT,TE,EE+lowE+lensing)

$$\Omega_{B0}h^2 = 0.02237 \pm 0.00015$$

$$\eta_{B0} \equiv \frac{n_{B0} - \bar{n}_{B0}}{n_{\gamma 0}} \simeq \frac{n_{B0}}{n_{\gamma 0}} \simeq 273.5\,\Omega_{B0}h^2 \times 10^{-10} = (6.12 \pm 0.04) \times 10^{-10}$$

- Consistent with (older) BBN determination but more precise and accurate
- Asymmetry coincides with matter abundance since there is no evidence of primordial antimatter…..not so far at least (see AMS-02 results and Poulin,Salati,Cholis,Kamionkowski,Silk 1808.08961)