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MAX-PLANCK-GESELLSCHAFT

The MSW potentials induced by ultralight mediators

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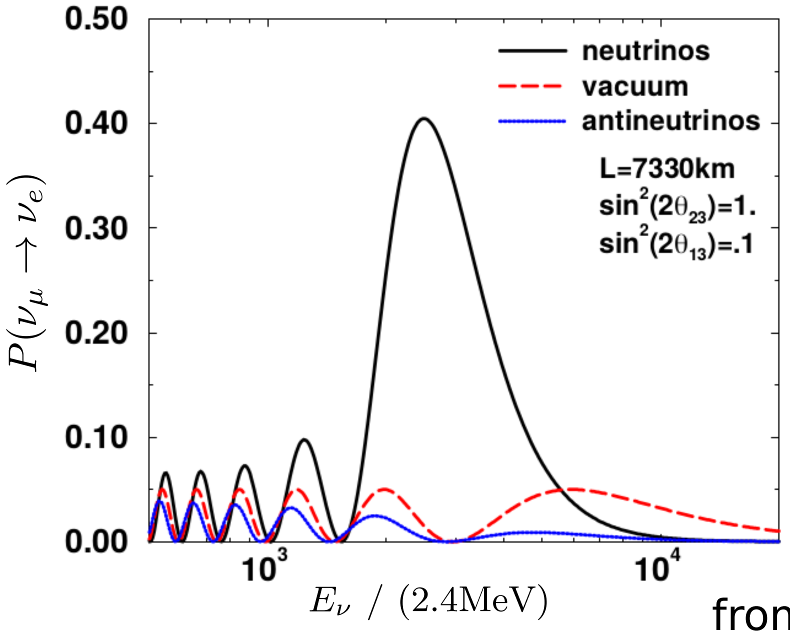
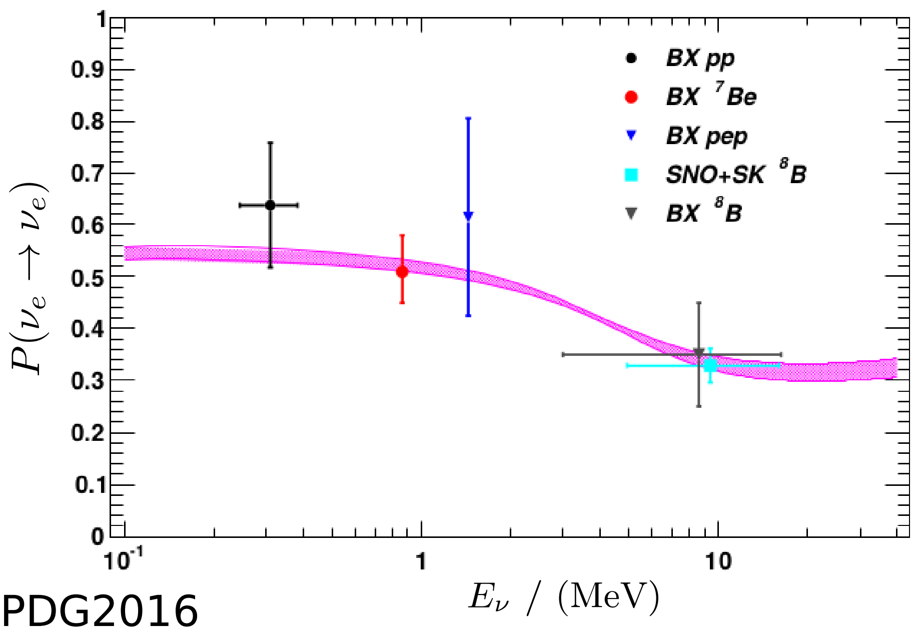
Talk based on: 1909.07505.

NEPLES2019 workshop, http://events.kias.re.kr/h/NEPLES2019/

The Mikheyev-Smirnov-Wolfenstein effect





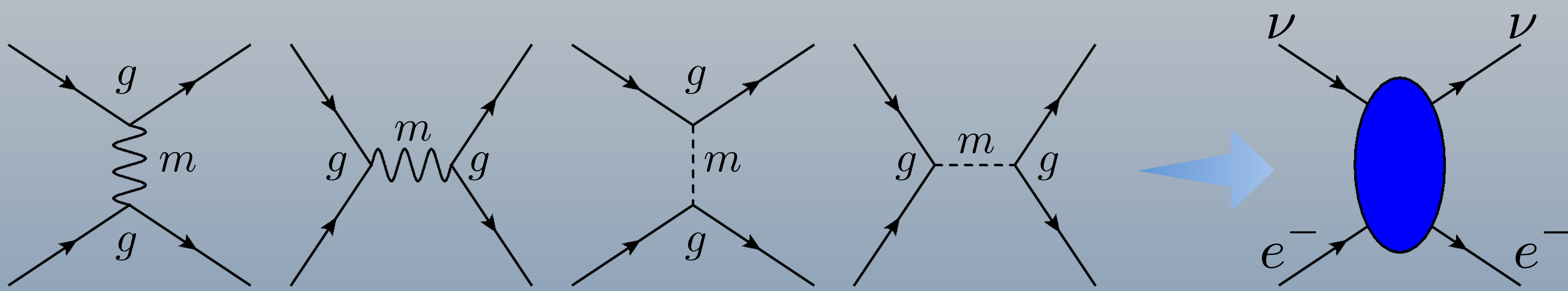
from PDG2016

$$i\frac{d}{dt}\begin{pmatrix}\nu_e\\ \nu_\mu\\ \nu_\tau\end{pmatrix} = H\begin{pmatrix}\nu_e\\ \nu_\mu\\ \nu_\tau\end{pmatrix}$$

The Wolfenstein potential:
$V = \sqrt{2}G_F n_e$

$$H = \frac{1}{2E}U\begin{pmatrix}m_1^2 & & \\ & m_2^2 & \\ & & m_3^2\end{pmatrix}U^\dagger + \begin{pmatrix}V & & \\ & 0 & \\ & & 0\end{pmatrix}$$

What if ν's have new interactions?



$$V = n_e \frac{g^2}{m^2}$$

Vector interactions: $H \to H + V$

Scalar interactions: $m_\nu \to m_\nu + V$

Comments on scalar interactions:

- Typically suppressed by m_ν/E_ν.

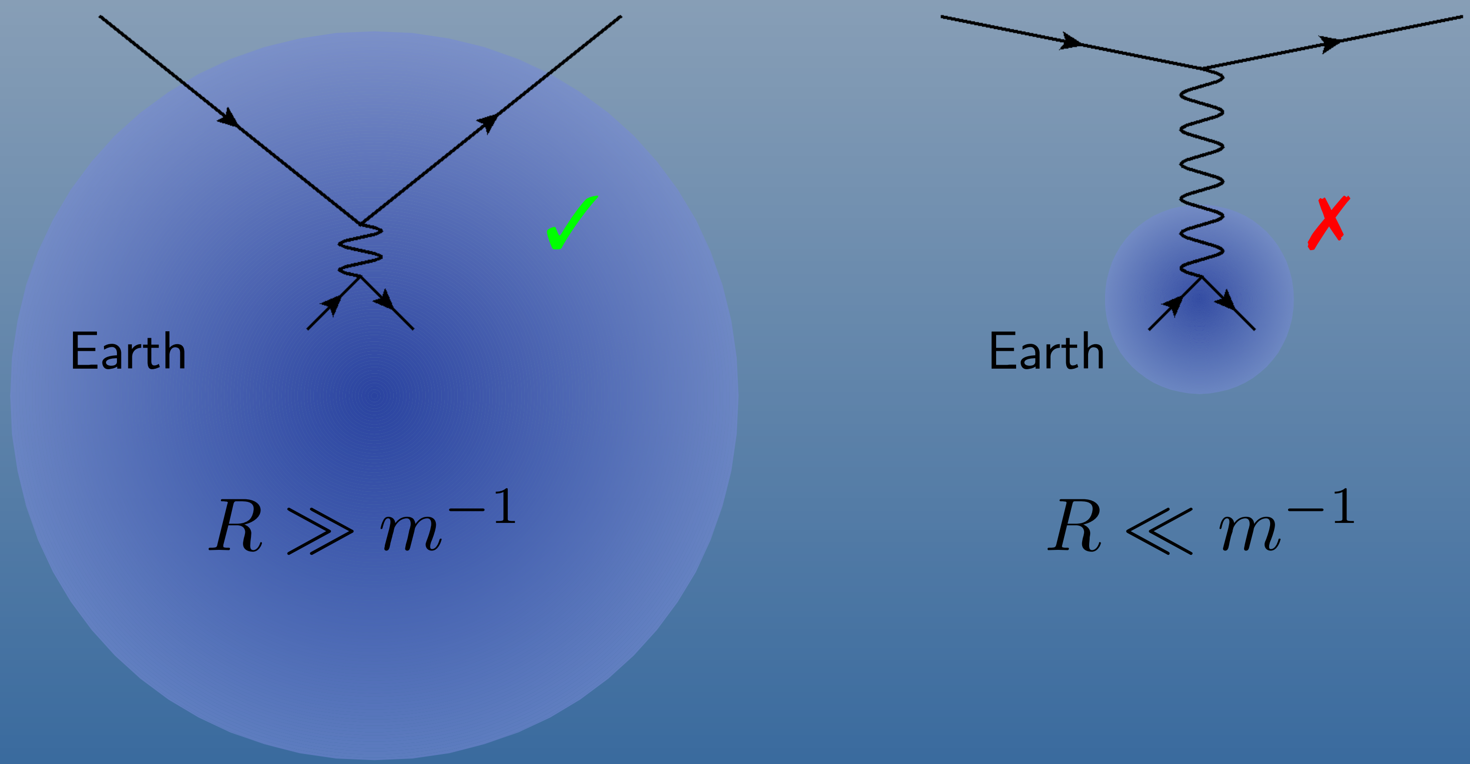
E.g. $m_\nu \sim 0.01$ eV, $E_\nu \sim 100$ MeV,
$\Rightarrow m_\nu/E_\nu \sim 10^{-10}$

- Impossible to have any significant effect in any realistic experiments.
- at Earth: $m_\nu \to m_\nu + 0.01$ eV $\Rightarrow$ in SN: $m_\nu \to m_\nu + 100$ GeV $\Rightarrow$ no SN ν!

$$V = n_e \frac{g^2}{m^2}$$

Q: always valid?

A: No!



Let's re-derive the MSW effect!

- Consider a simple system, only neutrino (ν) + electron (e) + vector boson (A^μ)
- Electrons are fixed by external forces, e.g., the electromagnetic (EM) force
- Neutrinos propagate through these electrons
- There are interactions between ν and e, mediated by the vector boson A^μ.

$$\mathcal{L} \supset \overline{\nu} i \partial\!\!\!/ \nu - m_\nu \overline{\nu}\nu - g\overline{\nu} A\!\!\!/ \nu - g\bar{e} A\!\!\!/ e - \frac{m^2}{2} A^\mu A_\mu$$

$\mathcal{L} \Rightarrow$EOM (equation of motion) $\Rightarrow$ MSW effect

EOM:

$$i\partial\!\!\!/ \nu - m_\nu \nu - g A\!\!\!/ \nu = 0,$$

$$\left[\partial^2 + m^2\right] A^\mu - g\overline{\nu}\gamma_\mu \nu - g\bar{e}\gamma_\mu e = 0$$

$$\cdots \text{ EOM of electrons} \cdots$$

Let's re-derive the MSW effect!

EOM (equation of motion)

$$i\partial\!\!\!/\nu - m_\nu \nu - g A\!\!\!/\nu = 0,$$

$$\left[\partial^2 + m^2\right] A^\mu - g\bar{\nu}\gamma_\mu\nu - g\bar{e}\gamma_\mu e = 0$$

$$\cdots \text{ EOM of electrons} \cdots$$

Next, solve EOM

the fields of ν, A^μ, e are determined by the solutions of EOM

e$\Rightarrow A^\mu \Rightarrow \nu$

- e field almost determined by the EM force

  assuming: $A^\mu \ll$ the EM force
- A^μ field almost determined by e field

  assuming: ν field $\ll e$ field
- ν field almost determined by A^μ field

  if ν have no other interactions

Let's re-derive the MSW effect!

EOM

$$i\partial\!\!\!/\nu - m_\nu \nu - gA\!\!\!/\nu = 0$$

$$\left[\partial^2 + m^2\right] A^\mu - g\bar{\nu}\gamma_\mu\nu - g\bar{e}\gamma_\mu e = 0$$

assuming: ν field $\ll$ e field

see Peskin&Schroeder's textbook

Physical meaning of $g\bar{e}\gamma_\mu e$:

- $J^i \equiv \bar{e}\gamma^i e$ is the electric current density
- $n_e \equiv \bar{e}\gamma^0 e$ is the electron number density

Usually electrons in matter are almost at rest $\Rightarrow J^i = 0$.
(average velocity $=0$)

So

$$\bar{e}\gamma^\mu e = (n_e,\ 0,\ 0,\ 0)$$

Let's re-derive the MSW effect!

EOM

$$i\partial\!\!\!/\nu - m_\nu \nu - gA\!\!\!/\nu = 0$$

$$\left[\partial^2 + m^2\right] A^\mu - g\bar{\nu}\gamma_\mu \nu - g\bar{e}\gamma_\mu e = 0$$

$$\left[\partial^2 + m^2\right] A^\mu - g\bar{e}\gamma_\mu e = 0$$

$\bar{e}\gamma^\mu e = (n_e,\ 0,\ 0,\ 0)$
Hence
$A^\mu = (A^0,\ 0,\ 0,\ 0)$

$$\left[\partial^2 + m^2\right] A^0 - gn_e = 0$$

$\partial_t A^0 = 0$

$$\left[-\nabla^2 + m^2\right] A^0 - gn_e = 0$$

Let's re-derive the MSW effect!

The master formula

$$\left[-\nabla^2 + m^2\right] A^0 = g n_e$$

- It's a differential equation.
- Given the distribution of electron density (n_e), we can always figure out A^0 by solving the equation.
- Physical meaning: $g n_e$ is the source that excites the A^0 field.
- $\left[-\nabla^2 + m^2\right] A^0 = g(n_{e1} + n_{e2} + n_{e3} + \cdots)$.

Let's re-derive the MSW effect!

EOM

$$i\partial\!\!\!/\nu - m_\nu \nu - g A\!\!\!/\nu = 0$$

$$\left[\partial^2 + m^2\right] A^\mu - g\bar{e}\gamma_\mu e = 0$$

$$\left[-\nabla^2 + m^2\right] A^0 = g n_e$$

Given n_e, determine $A^\mu = (A^0,\ 0,\ 0,\ 0)$

Given A^μ, how does ν propagate?

$(i\partial\!\!\!/ - g A\!\!\!/)\nu - m_\nu \nu = 0$

In the momentum space, $i\partial\!\!\!/ \to -p^\mu\gamma_\mu$ so:

$$p^\mu \to p^\mu + g A^\mu$$

$$E = p^0 \to E + V,\ V \equiv g A^0$$

In case of scalar: $m_\nu \to m_\nu + g\phi$

Let's re-derive the MSW effect!

Solve the differential equation

$$\left[-\nabla^2 + m^2\right] A^0 = g n_e$$

Spherical cows in a vacuum

There's a dairy farm. One day, the milk production was low. So the farmer called a physicist to help. The physicist then did a lot of calculations, and he said: "um, I have a solution, but it only works with spherical cows in a vacuum."

From Wikipedia: https://en.wikipedia.org/wiki/Spherical_cow

Let's re-derive the MSW effect!

Solve the differential equation

$$\left[-\nabla^2 + m^2\right] A^0 = g n_e$$

Let's assume a spherical cow:

$$n_e(r) = \begin{cases} 0 & (\text{for } r > R) \\ n_e & (\text{for } r \leq R) \end{cases},$$

the solution is

$$A^0(r) = \frac{g n_e}{m^2} F(r),$$

$$F(r) = \begin{cases} 1 - \frac{mR+1}{mr} e^{-mR} \sinh(mr) & (r \leq R) \\ \frac{e^{-mr}}{mr} \left[mR \cosh(mR) - \sinh(mR)\right] & (r > R) \end{cases}$$

Interesting limits of the spherical cow solution

- very compact cow ($R \to 0$ while $N_e \equiv \frac{4}{3}\pi R^3 n_e = \mathrm{const.}$):

$$V(r) = \frac{g^2}{m^2} N_e \frac{e^{-mr}}{4\pi r}$$

$\Rightarrow$ the well-known Yukawa potential

- long-range interacting cow ($m \to 0$):

$$V(r) = g^2 n_e \times \begin{cases} \frac{3R^2 - r^2}{6} & (r \leq R) \\ \frac{R^3}{3}\frac{1}{r} & (r > R) \end{cases}$$

$\Rightarrow$ the well-known Coulomb potential

Interesting limits of the spherical cow solution

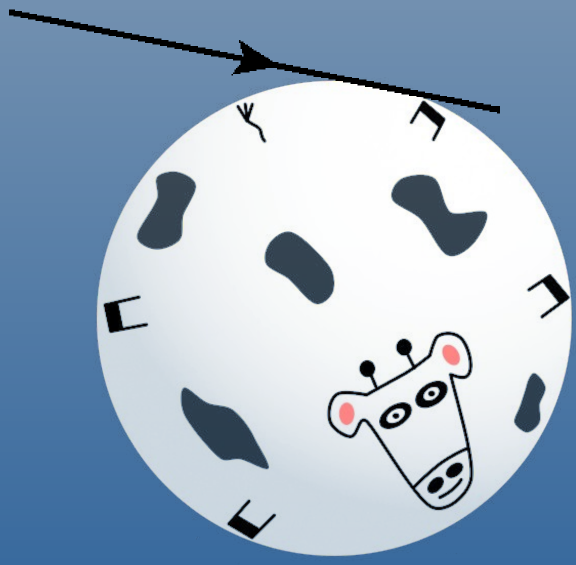
- contact-interacting cow ($m \to \infty$):

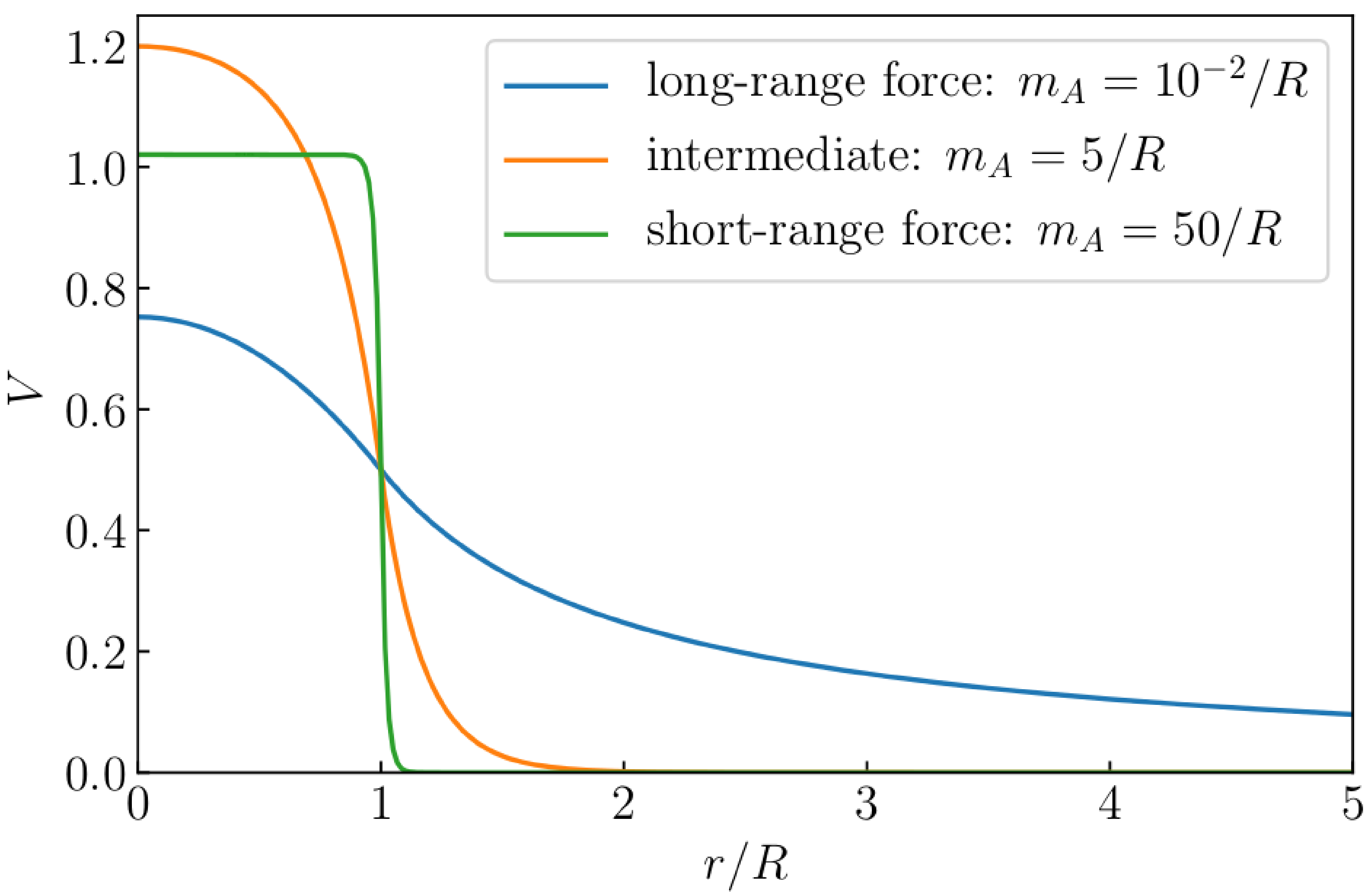
$$V(r) = \begin{cases} \frac{g^2}{m^2} n_e & (r \leq R) \\ 0 & (r > R) \end{cases}$$

$\Rightarrow$ the standard Wolfenstein potential

- ν at cow's skin ($|R - r| \ll m^{-1} \ll R$):

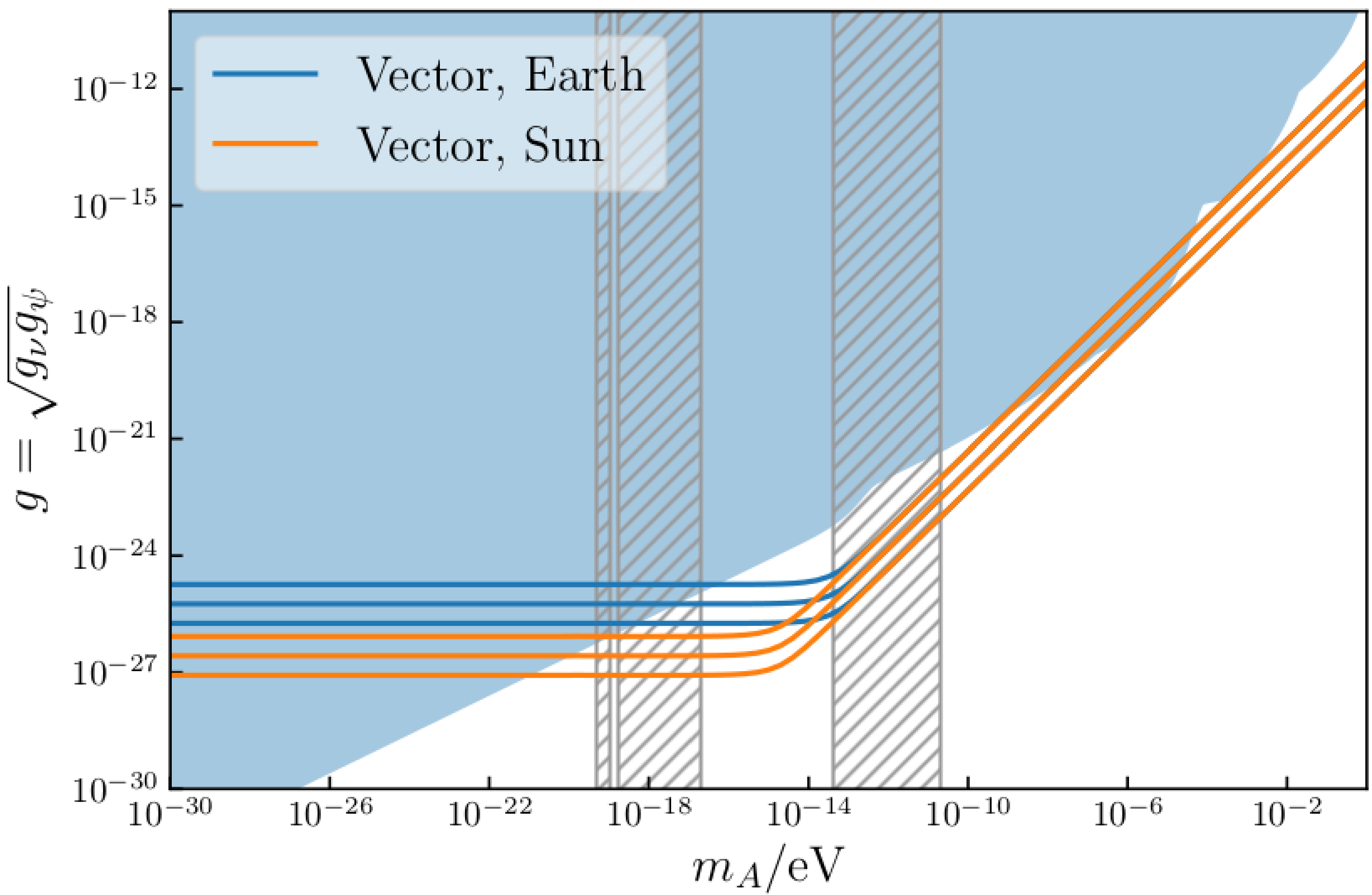
$$V(r) = \frac{1}{2} \frac{g^2}{m^2} n_e$$



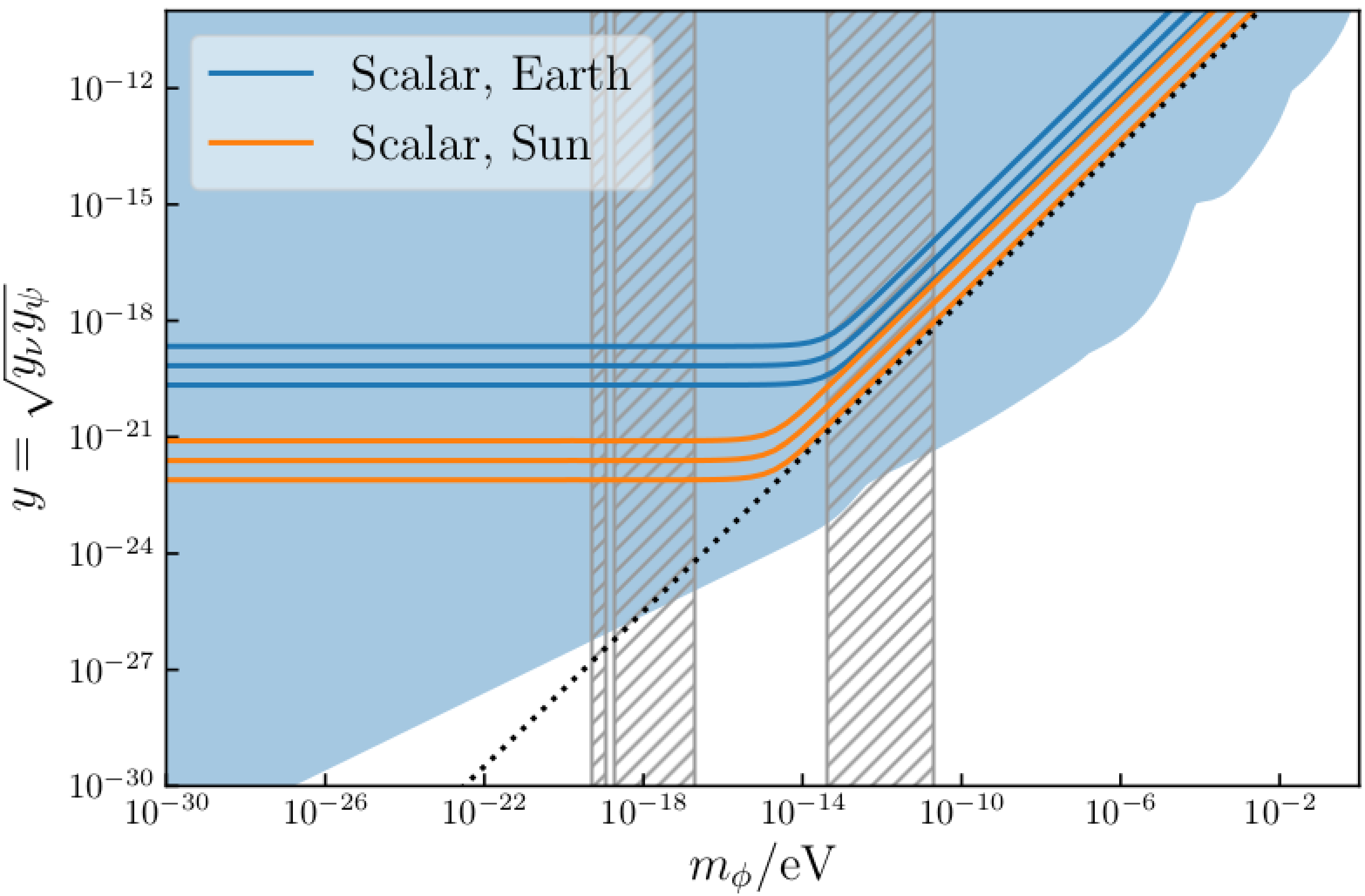


long-range force: $m_A = 10^{-2}/R$
intermediate: $m_A = 5/R$
short-range force: $m_A = 50/R$
V
1.2
1.0
0.8
0.6
0.4
0.2
0.0
0
1
2
3
4
5
r/R

Vector effects, assuming $\frac{V}{V_{\rm SM}} = \{10^{-2},\ 10^{-1},\ 1\}$



Scalar effects, assuming $\delta m_\nu = \{10^{-3},\ 10^{-2},\ 10^{-1}\}$ eV



Summary

Take-home messages

$$\left[-\nabla^2 + m^2\right] A^0 = g n_e$$

1. This is a powerful cow

   Spherical cow in a vacuum
   $\Rightarrow$ Yukawa, Coulomb, Wolfenstein potentials.

2. Scalar matter effect doesn't work

   requires $10^{10} G_F$,
   excluded by, e.g., ν free streaming, $m_\nu \to m_\nu + \delta m$ in supernova.

3. Vector matter effect works

   High energy: normal NSI
   Low energy: e.g. $m_A \in [2 \times 10^{-17}, 4 \times 10^{-14}]$ eV, $g \sim 10^{-25}$.

Backup

Vector $\rightarrow$ Scalar

$$\mathcal{L} \supset \overline{\nu} i \partial\!\!\!/ \nu - m_\nu \overline{\nu}\nu - g\overline{\nu} A\!\!\!/ \nu - g\bar{e} A\!\!\!/ e - \frac{m^2}{2} A^\mu A_\mu$$

$$i\partial\!\!\!/ \nu - m_\nu \nu - g A\!\!\!/ \nu = 0$$

$$\left[\partial^2 + m^2\right] A^\mu - g\bar{e}\gamma_\mu e = 0$$

$$\mathcal{L} \supset \overline{\nu} i \partial\!\!\!/ \nu - m_\nu \overline{\nu}\nu - g\overline{\nu}\phi\nu - g\bar{e}\phi e - \frac{m^2}{2}\phi^2$$

$$i\partial\!\!\!/ \nu - m_\nu \nu - g\phi\nu = 0$$

$$\left[\partial^2 + m^2\right]\phi - g\bar{e}e = 0$$

Vector $\to$ Scalar

$$\left[\partial^2 + m^2\right] A^\mu - g\bar{e}\gamma_\mu e = 0 \quad \Rightarrow \quad \left[\partial^2 + m^2\right] \phi - g\bar{e}e = 0$$

$$\left[-\nabla^2 + m^2\right] A^0 - g n_e = 0 \quad \Rightarrow \quad \left[-\nabla^2 + m^2\right] \phi - g n_e = 0$$

Same equation $\Rightarrow$ same solution!

But!

$$i\not{\partial}\nu - m_\nu \nu - g\not{A}\nu = 0 \quad \Rightarrow \quad i\not{\partial}\nu - m_\nu \nu - g\phi\nu = 0$$

$$p^\mu \to p^\mu + gA^\mu$$

$$E \to E + gA^0$$

$$\Rightarrow \quad m_\nu \to m_\nu + g\phi$$