

Modeling superfluidity in neutron stars with Brussels-Montreal functionals

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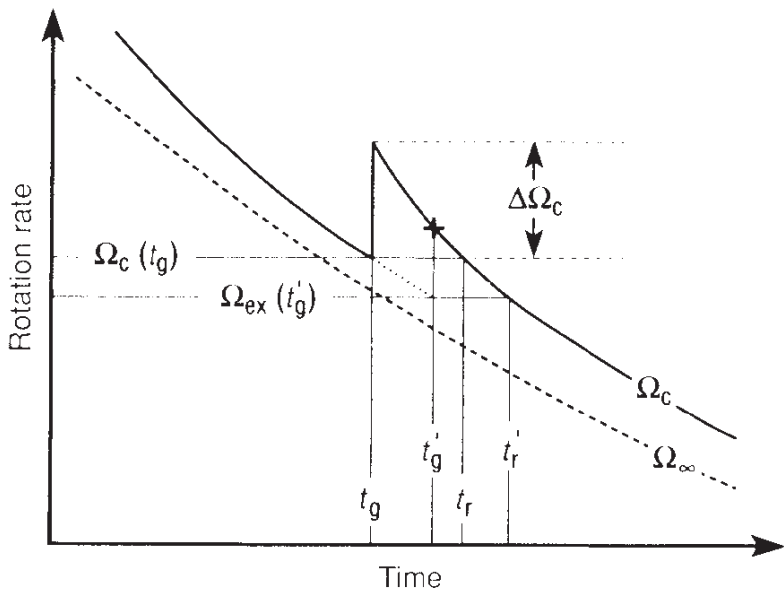
**Faculty
of Physics**

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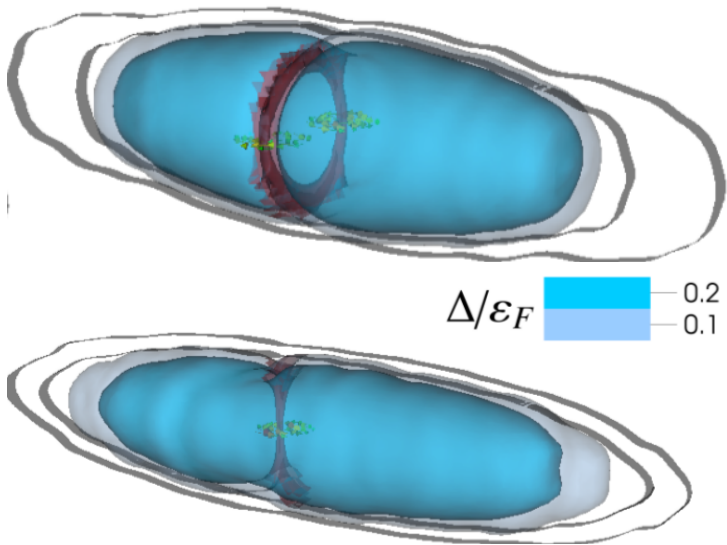


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28th February 2020, Karpacz



Link et al., Nature 359, 616 (1992)



Wlazłowski et al., Phys. Rev. Lett. 120, 253002 (2018)

$$\rho(r) = \sum_k |v_k(r)|^2 \quad \tau(r) = \sum_k |\nabla v_k(r)|^2 \quad \nu(r) = \sum_k u_k(r)v_k^*(r)$$

$$h(r) \equiv h(\rho(r), \tau(r), \nabla\rho(r), \vec{j}(r))$$

$$\Delta(r) \equiv \Delta(\rho(r), \nabla\rho(r), \nu(r))$$

Bogolubov-de Gennes equations

$$\begin{pmatrix} h(r) & \Delta(r) \\ \Delta^*(r) & -h^*(r) \end{pmatrix} \begin{pmatrix} u_k(r) \\ v_k(r) \end{pmatrix} = \epsilon_k \begin{pmatrix} u_k(r) \\ v_k(r) \end{pmatrix}$$



$$\varepsilon(\rho, \vec{\nabla}\rho, \nu, \tau, \mathbf{j}) = \frac{\hbar^2}{2M}\tau + \varepsilon_\rho(\rho) + \varepsilon_\tau(\rho, \tau, \mathbf{j}) + \varepsilon_{\Delta\rho}(\rho, \vec{\nabla}\rho) + \varepsilon_\pi(\rho, \vec{\nabla}\rho, \nu)$$

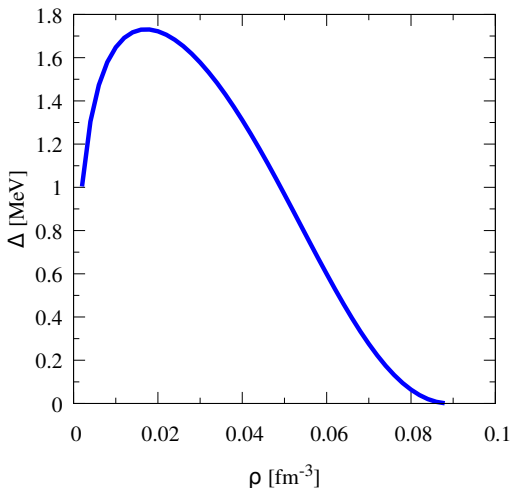
$$h(r) = \frac{\delta\varepsilon}{\delta\rho} - \nabla \frac{\delta\varepsilon}{\delta\tau} \nabla - \frac{i}{2} \left\{ \frac{\delta\varepsilon}{\delta\mathbf{j}}, \nabla \right\}$$

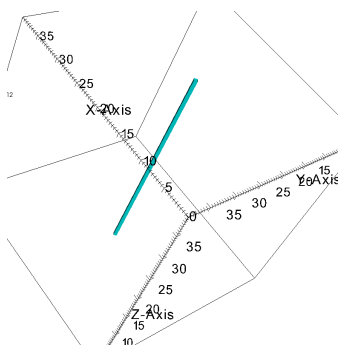
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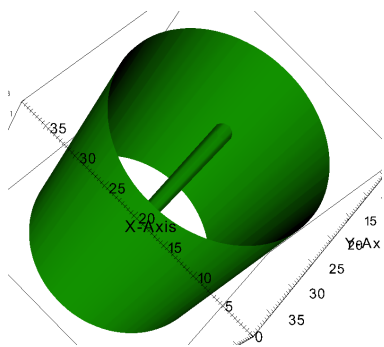
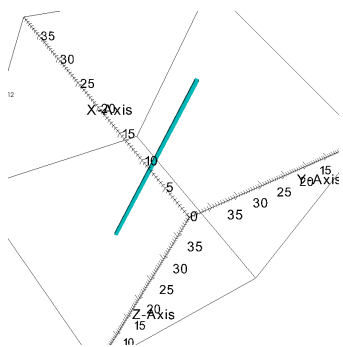
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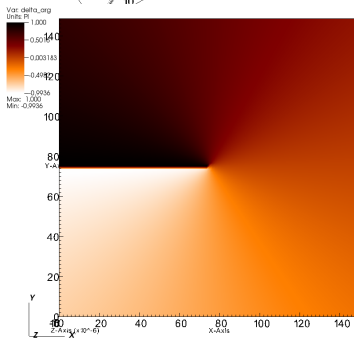
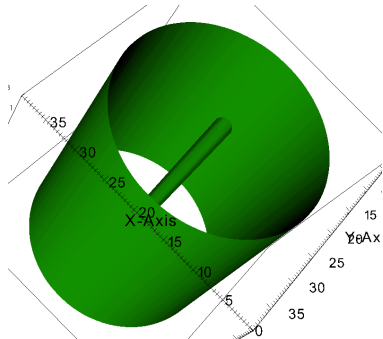
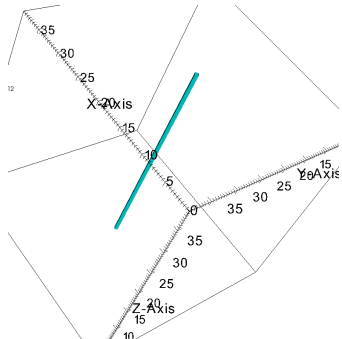
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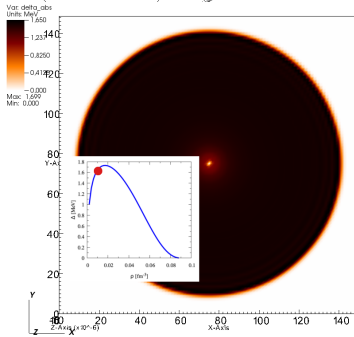
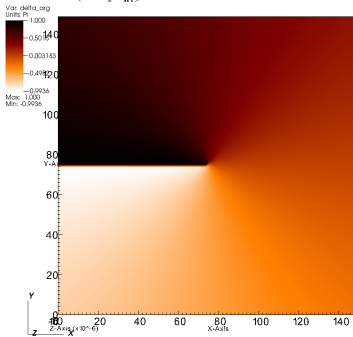
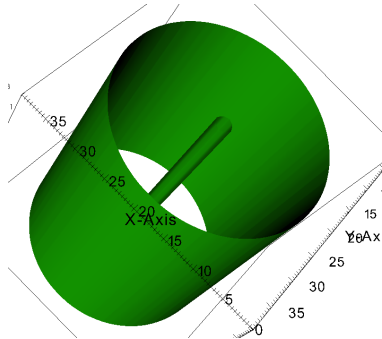
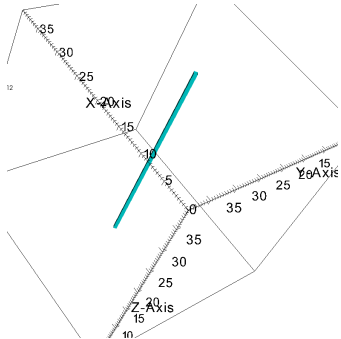
$$\Delta(r) = \frac{\delta\varepsilon}{\delta\nu}$$

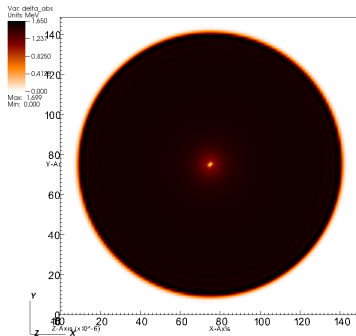


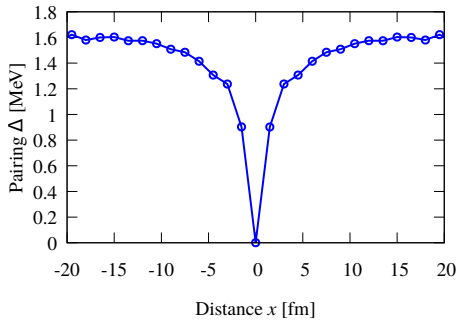
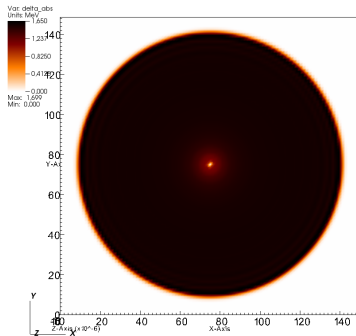


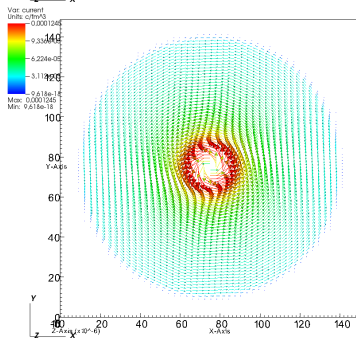
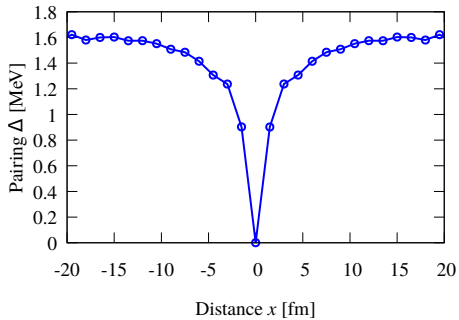
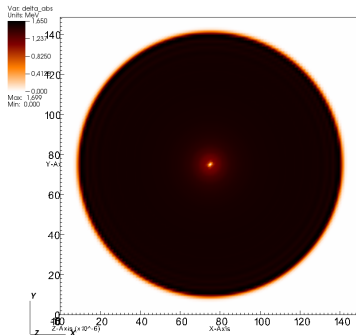


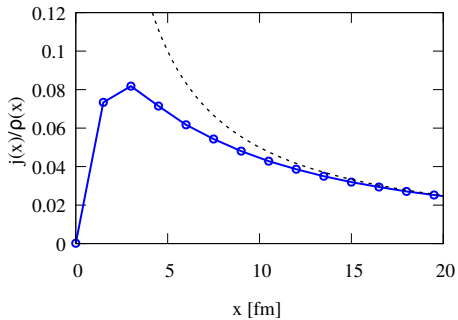
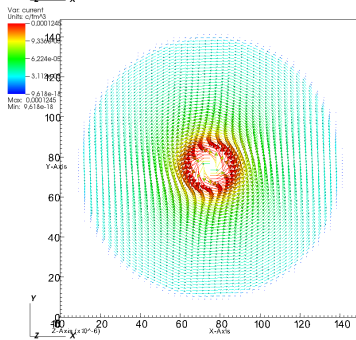
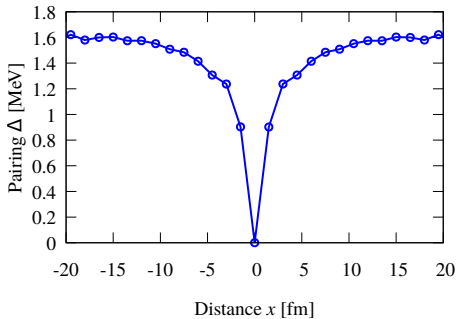
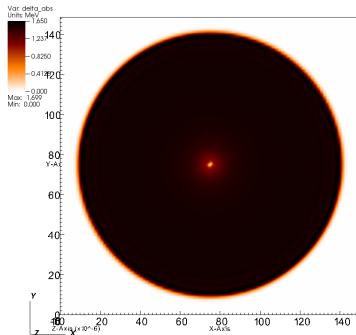


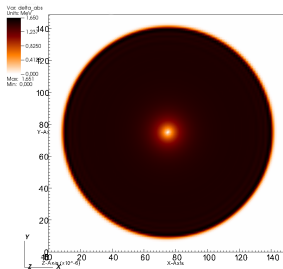
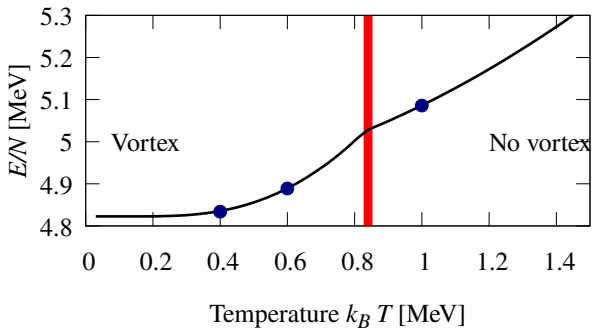


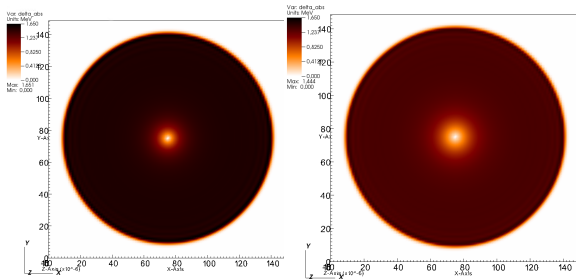
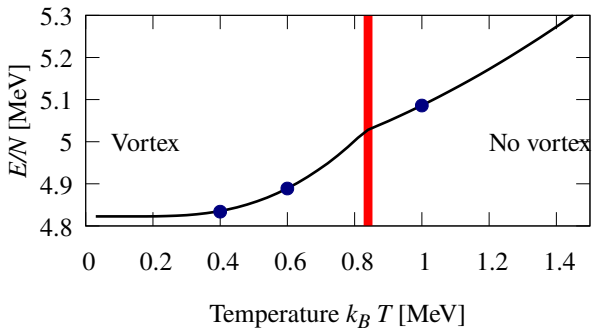


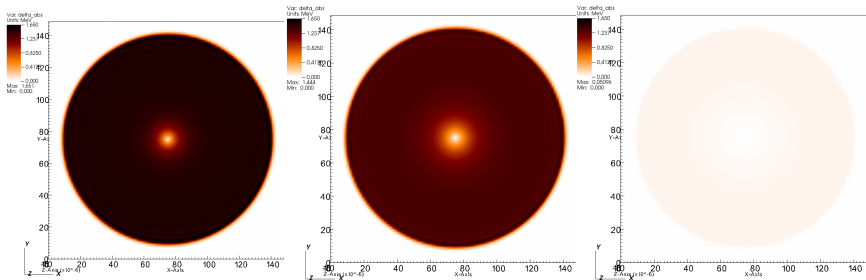
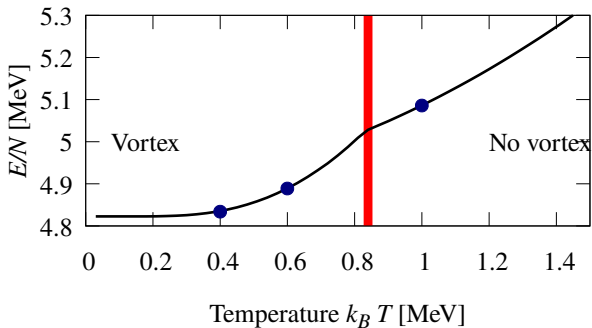




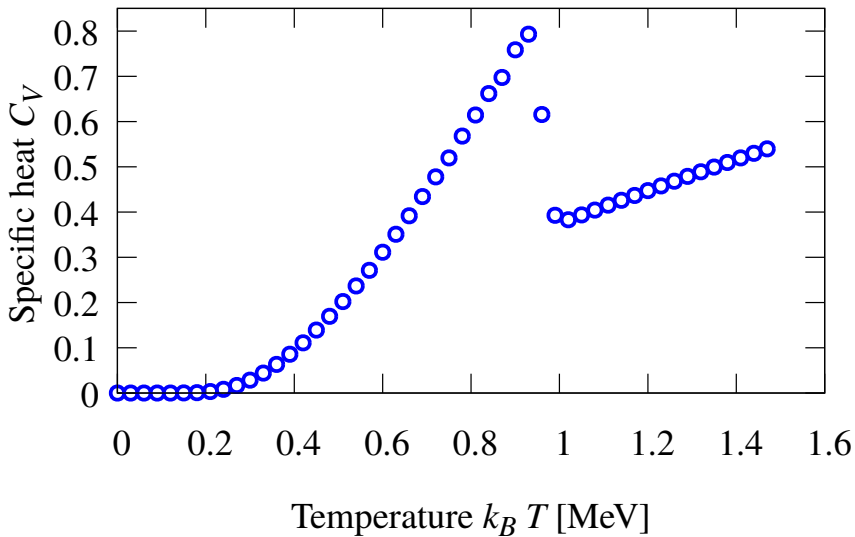








$$C_V = \frac{\partial E}{\partial T}$$



Outlook

- ★ heat capacity
- ★ full 3D dynamics
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Take-home message

- ★ BSk fitted for both regimes: neutron matter and nuclei (clusters & vortices)
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Thank you!