

Co-SIMP Miracle

Juri Smirnov



PHENO 2020: 05/05/20

In collaboration with: John Beacom (OSU, CCAPP)

Dark Matter is a New Particle

Dark Matter is a New Particle

Not ordinary Matter:

$$\Omega_{\text{DM}} \gg \Omega_{\text{Baryons}}$$

(CMB) (BBN, CMB)

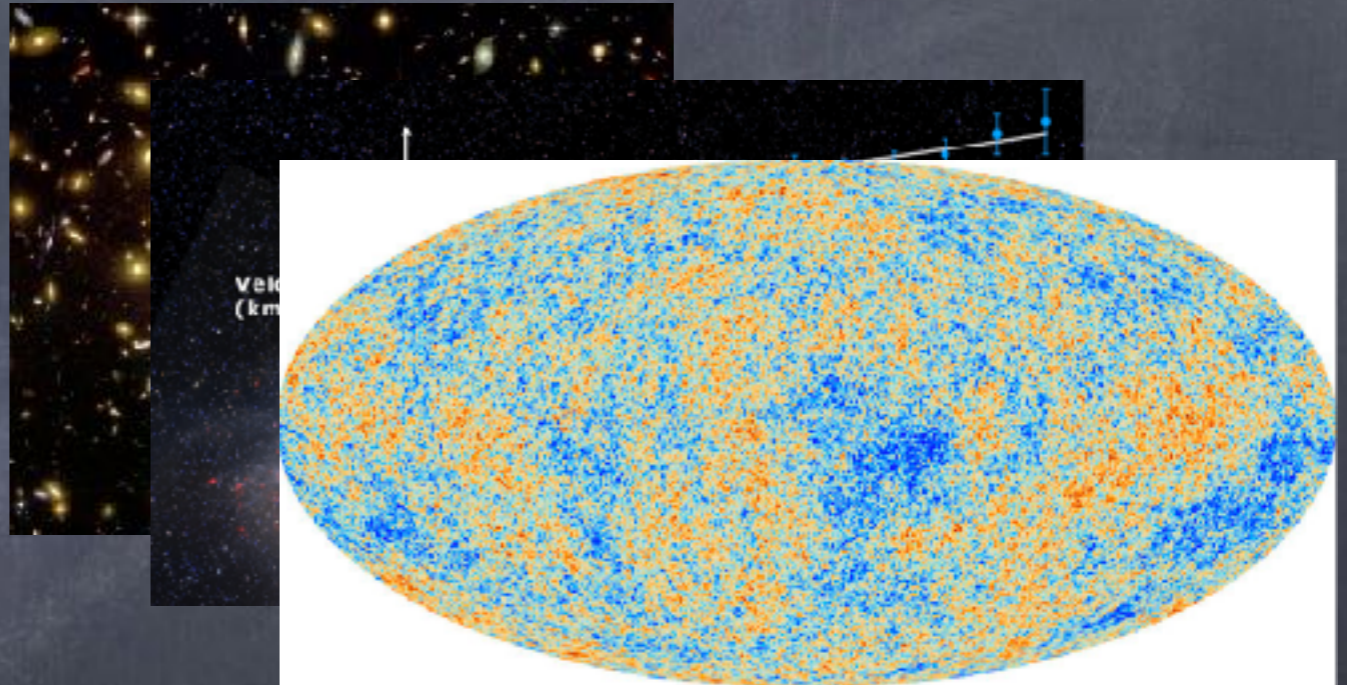
Dark Matter is a New Particle

Not ordinary Matter:

$$\Omega_{\text{DM}} \gg \Omega_{\text{Baryons}}$$

(CMB) (BBN, CMB)

Not MOND:



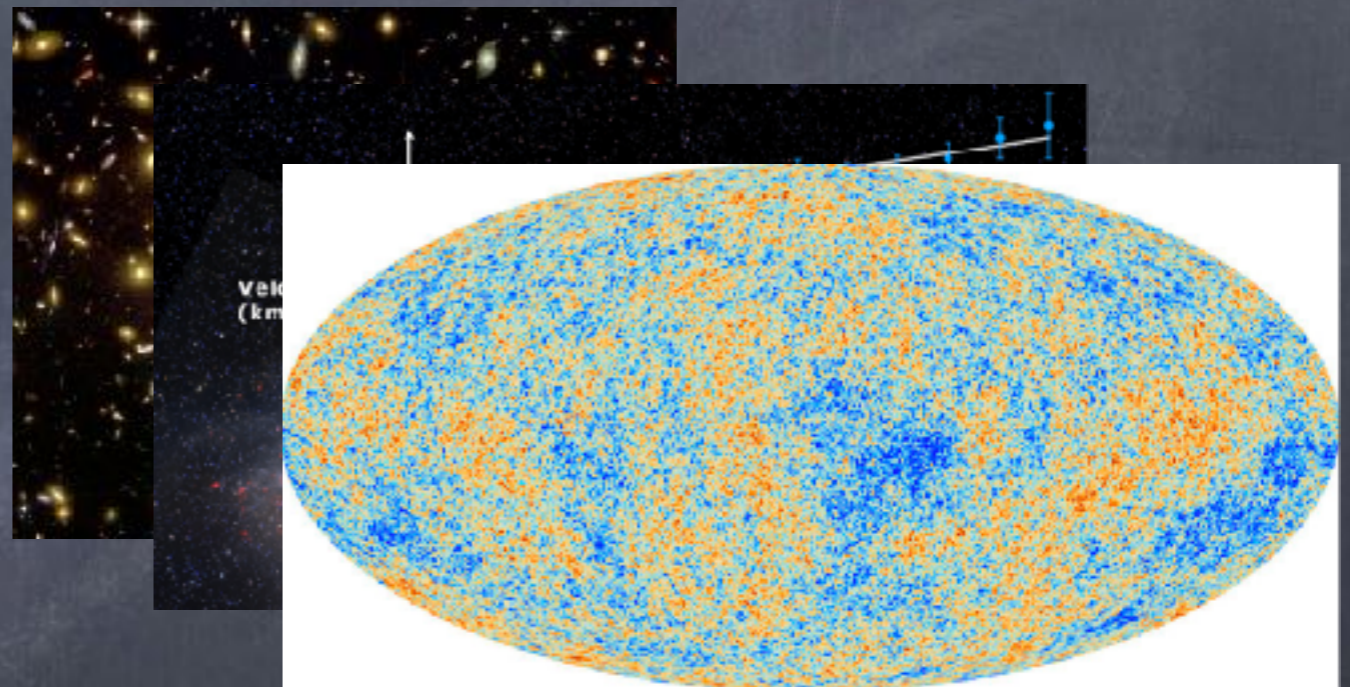
Dark Matter is a New Particle

Not ordinary Matter:

$$\Omega_{\text{DM}} \gg \Omega_{\text{Baryons}}$$

(CMB) (BBN, CMB)

Not MOND:

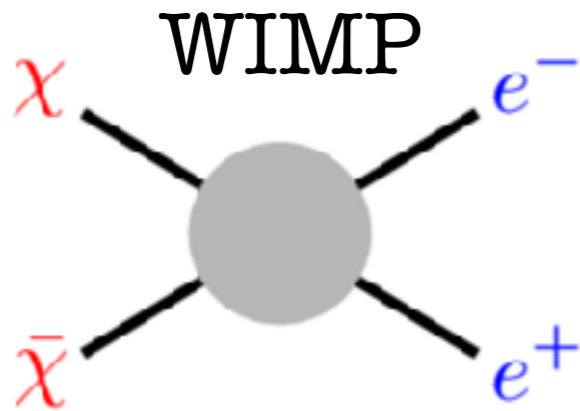


Not light Neutrinos:

$$\Omega_{\nu} \approx 0.02 \left(\frac{m_{\nu}}{\text{eV}} \right)$$

Production

Freezeout

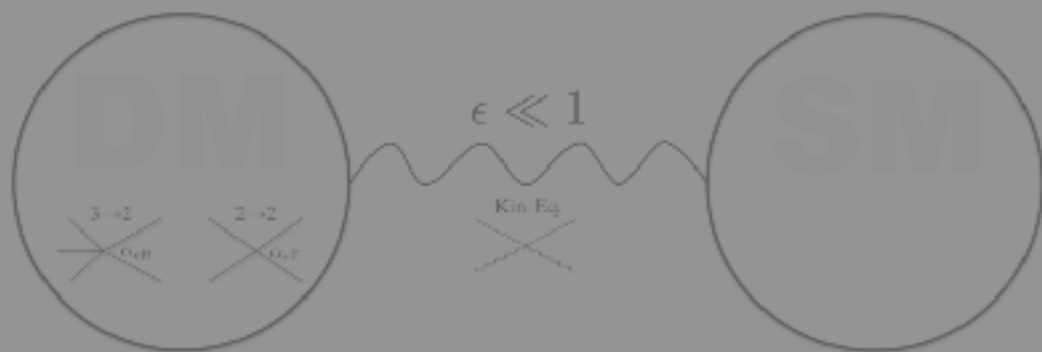


Zeldovich, Lee, Weinberg, Steigman, Turner,...

$$\Gamma_{\text{DM}} = \langle \sigma v_{\text{rel.}} \rangle n_{\text{DM}} > H(T)$$

$$\Omega_{\text{DM}} h^2 \approx \frac{0.12}{\langle \sigma v_{\text{rel.}} \rangle [25 \text{TeV}]^2}$$

SIMP

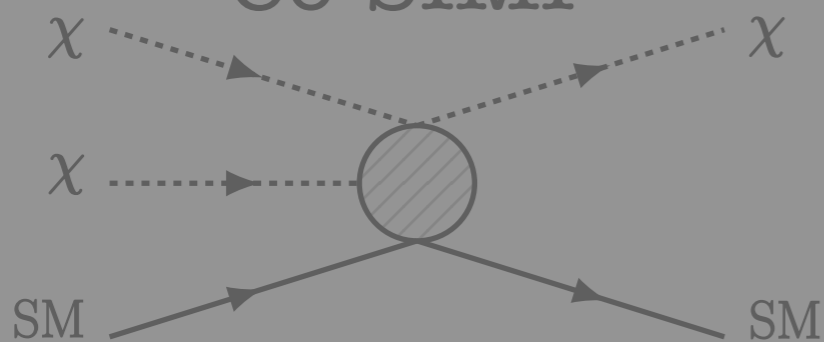


Hochberg, Kuflik, Volansky, Wacker

$$\Gamma_{\text{DM}} = \langle \sigma_{3 \rightarrow 2} v_{\text{rel.}}^2 \rangle n_{\text{DM}}^2 > H(T)$$

$$\Omega_{\text{DM}} h^2 \approx \left(\frac{\text{MeV}}{m_{\text{DM}}} \right) \frac{0.12}{\sqrt{\langle \sigma_{3 \rightarrow 2} v_{\text{rel.}}^2 \rangle [3 \text{MeV}]^5}}$$

Co-SIMP

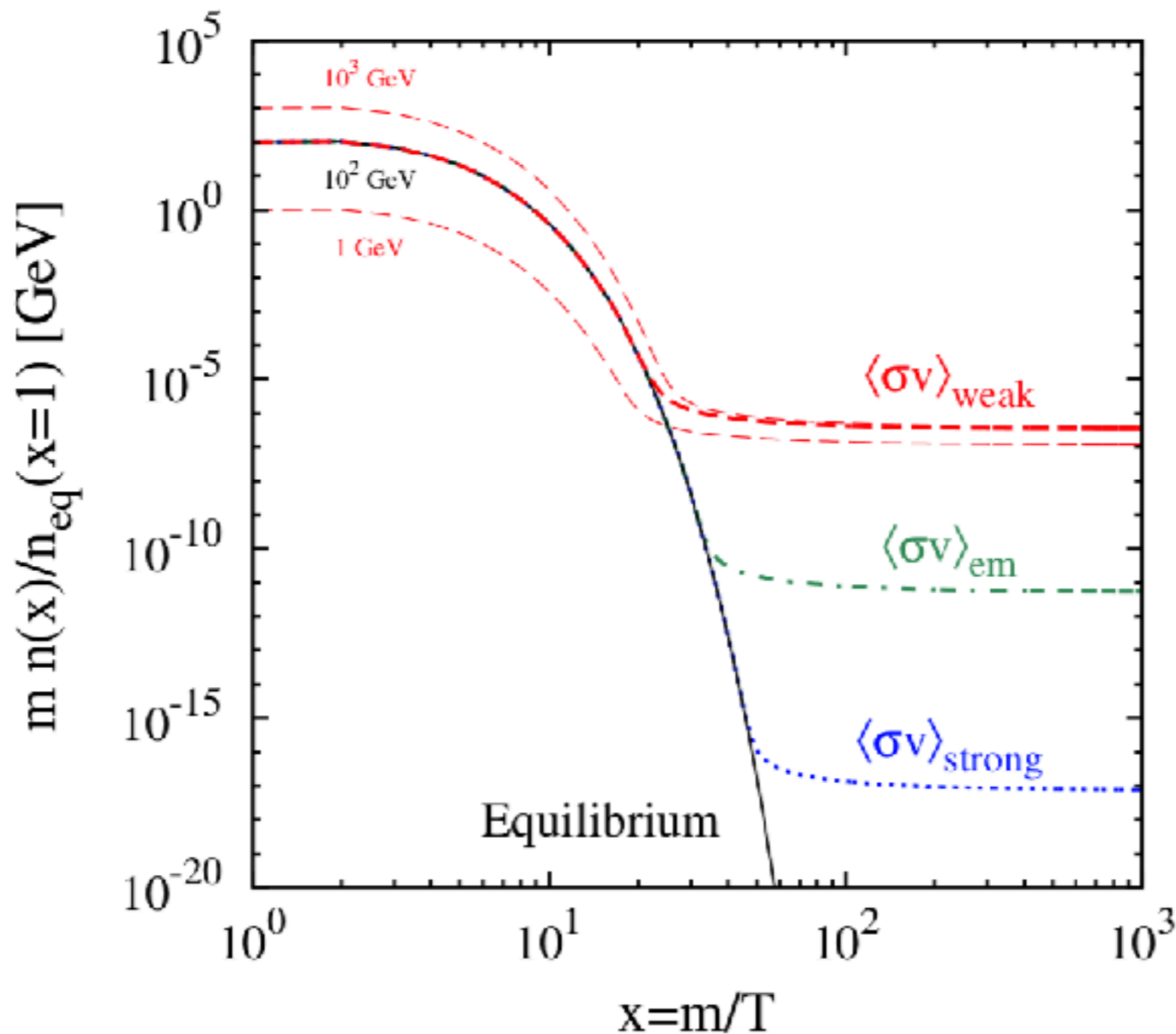


Smirnov, Beacom

$$\Gamma_{\text{DM}} = \langle \sigma_{3 \rightarrow 2} v_{\text{rel.}}^2 \rangle n_{\text{DM}} n_{\text{SM}} > H(T)$$

$$\Omega_{\text{DM}} h^2 \approx \left(\frac{m_{\text{DM}}}{\text{GeV}} \right)^3 \frac{0.12}{\langle \sigma_{3 \rightarrow 2} v_{\text{rel.}}^2 \rangle [3 \text{GeV}]^5}$$

WIMP Freezeout



Steigman, Dasgupta, Beacom
arXiv: 1204.3622

Cartoon Overview of Approach

Dark Matter Annihilation $\langle \sigma v_{\text{rel.}} \rangle$



Dark Matter Mass

Cartoon Overview of Approach

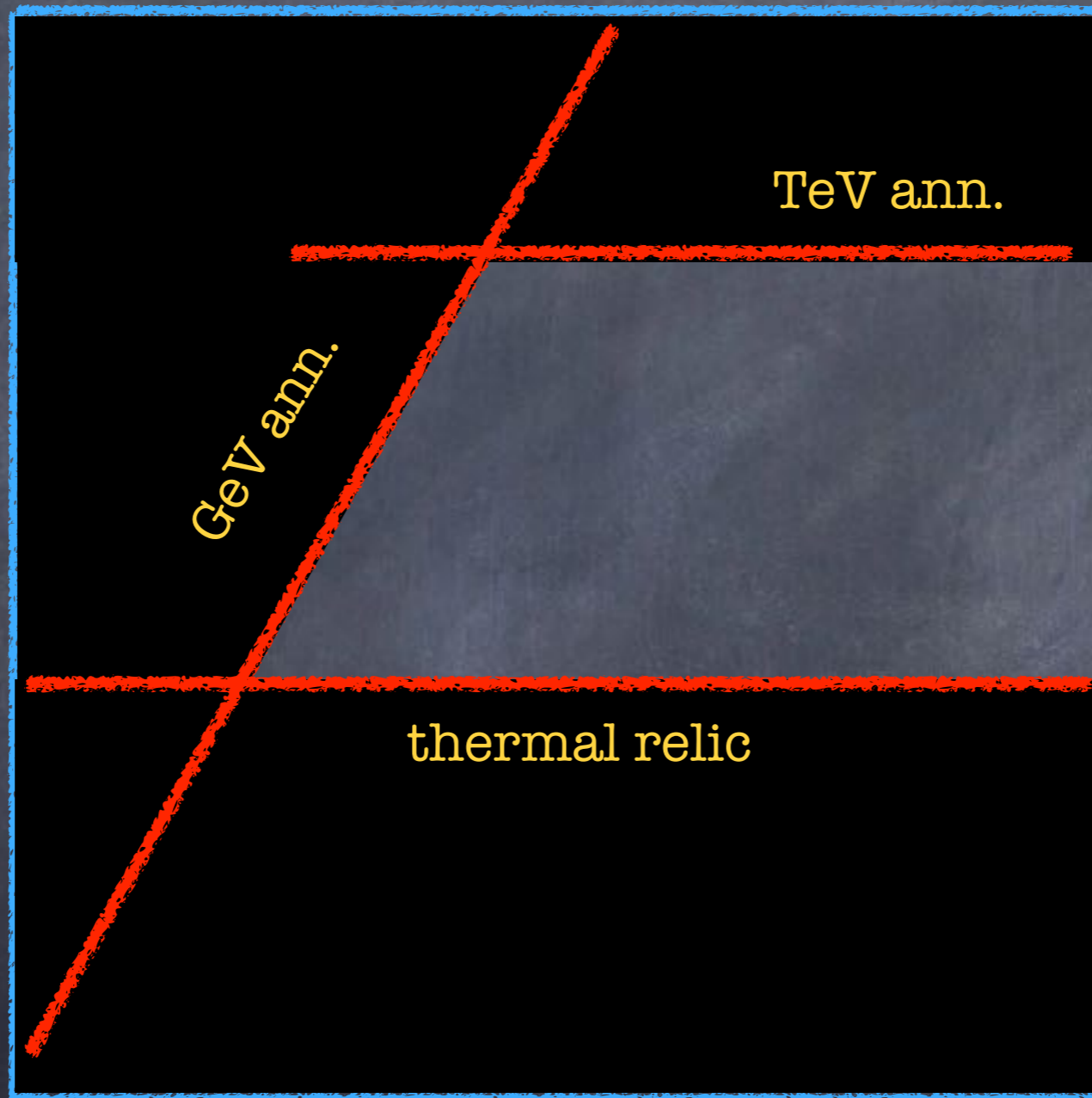
Dark Matter Annihilation $\langle \sigma v_{\text{rel.}} \rangle$



Dark Matter Mass

Cartoon Overview of Approach

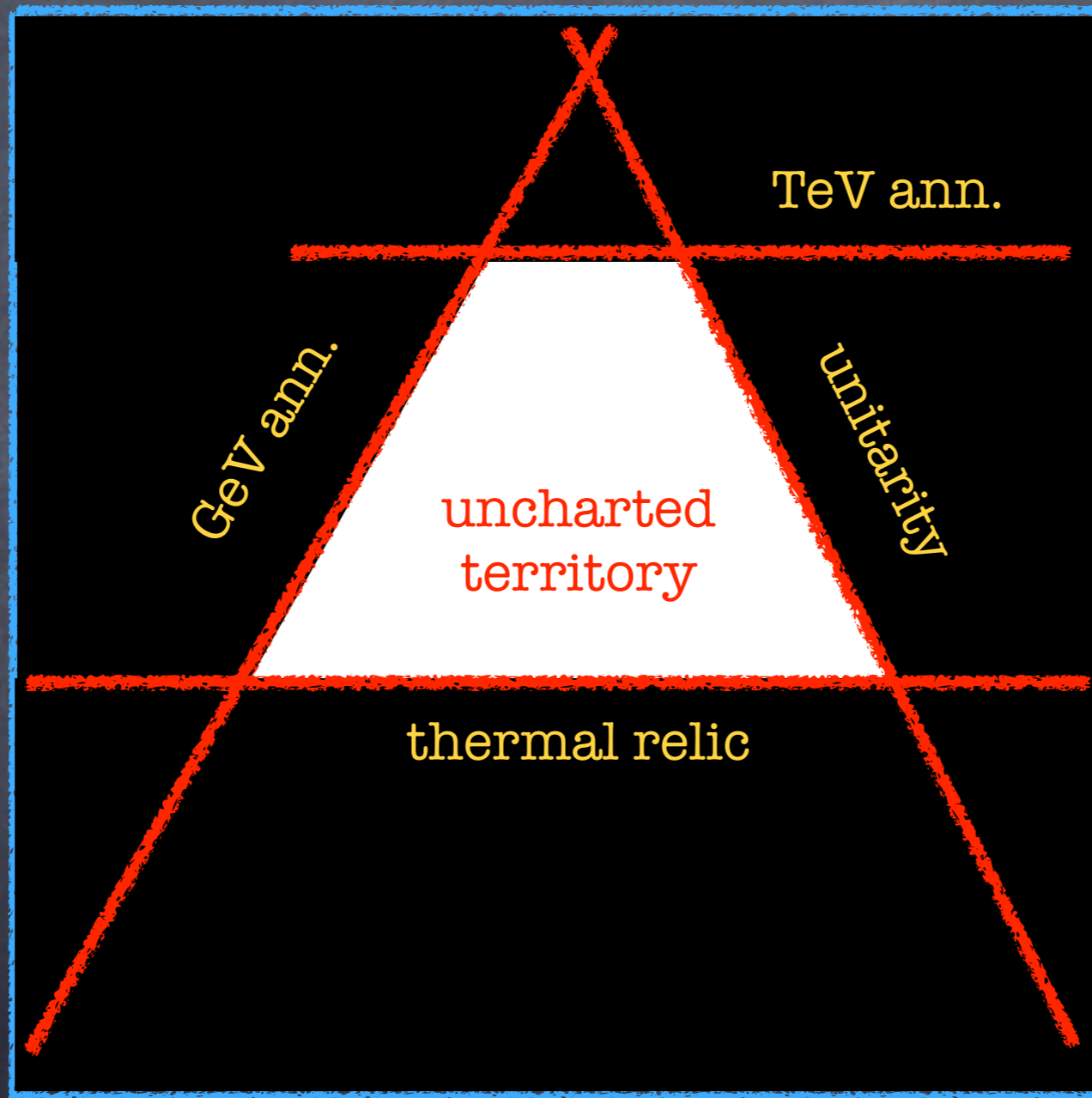
Dark Matter Annihilation $\langle\sigma v_{\text{rel.}}\rangle$



Dark Matter Mass

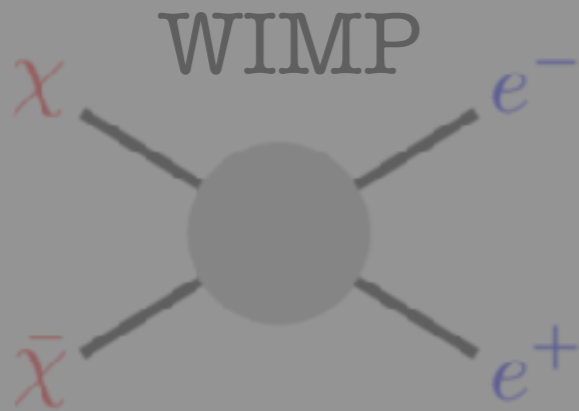
Cartoon Overview of Approach

Dark Matter Annihilation $\langle\sigma v_{\text{rel.}}\rangle$



Dark Matter Mass

Other Types of Freezeout

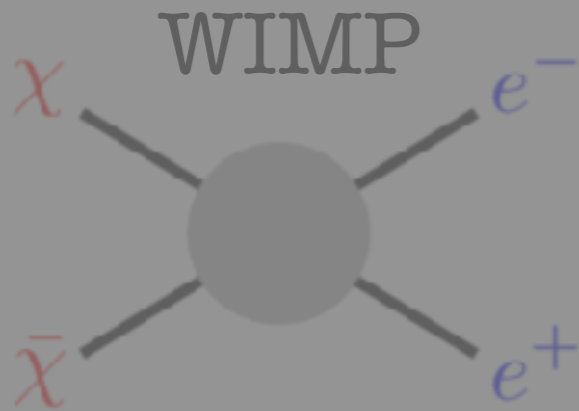


Zeldovich, Lee, Weinberg, Steigman, Turner,...

$$\Gamma_{\text{DM}} = \langle \sigma v_{\text{rel.}} \rangle n_{\text{DM}} > H(T)$$

$$\Omega_{\text{DM}} h^2 \approx \frac{0.12}{\langle \sigma v_{\text{rel.}} \rangle [25 \text{TeV}]^2}$$

Other Types of Freezeout

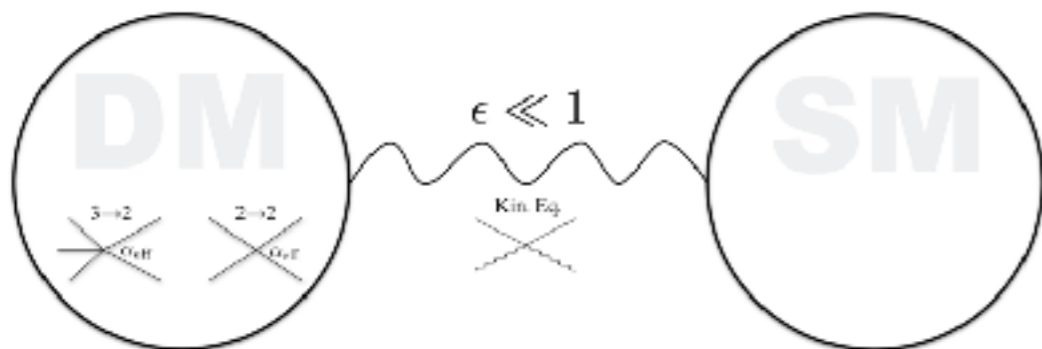


Zeldovich, Lee, Weinberg, Steigman, Turner,...

$$\Gamma_{\text{DM}} = \langle \sigma v_{\text{rel.}} \rangle n_{\text{DM}} > H(T)$$

$$\Omega_{\text{DM}} h^2 \approx \frac{0.12}{\langle \sigma v_{\text{rel.}} \rangle [25 \text{ TeV}]^2}$$

SIMP

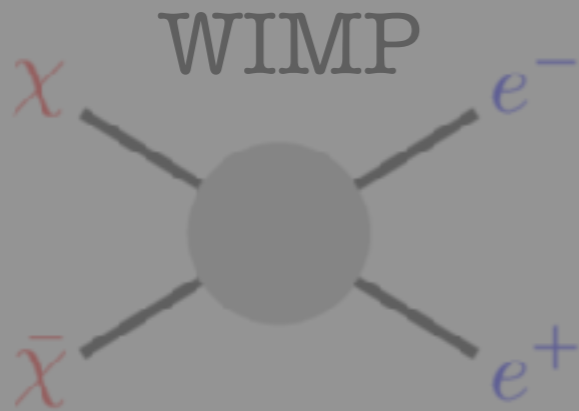


Hochberg, Kuflik, Volansky, Wacker

$$\Gamma_{\text{DM}} = \langle \sigma_{3 \rightarrow 2} v_{\text{rel.}}^2 \rangle n_{\text{DM}}^2 > H(T)$$

$$\Omega_{\text{DM}} h^2 \approx \left(\frac{\text{MeV}}{m_{\text{DM}}} \right) \frac{0.12}{\sqrt{\langle \sigma_{3 \rightarrow 2} v_{\text{rel.}}^2 \rangle [3 \text{ MeV}]^5}}$$

Other Types of Freezeout

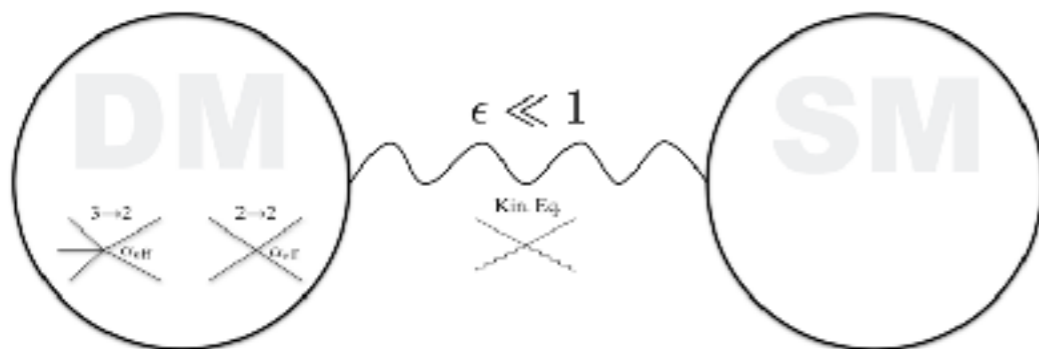


Zeldovich, Lee, Weinberg, Steigman, Turner,...

$$\Gamma_{\text{DM}} = \langle \sigma v_{\text{rel.}} \rangle n_{\text{DM}} > H(T)$$

$$\Omega_{\text{DM}} h^2 \approx \frac{0.12}{\langle \sigma v_{\text{rel.}} \rangle [25 \text{ TeV}]^2}$$

SIMP

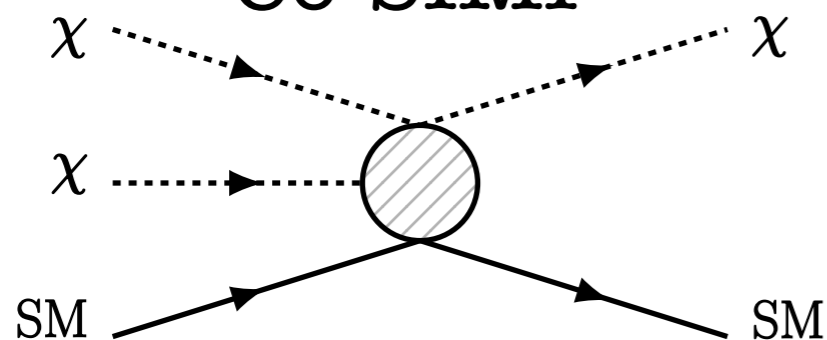


Hochberg, Kuflik, Volansky, Wacker

$$\Gamma_{\text{DM}} = \langle \sigma_{3 \rightarrow 2} v_{\text{rel.}}^2 \rangle n_{\text{DM}}^2 > H(T)$$

$$\Omega_{\text{DM}} h^2 \approx \left(\frac{\text{MeV}}{m_{\text{DM}}} \right) \frac{0.12}{\sqrt{\langle \sigma_{3 \rightarrow 2} v_{\text{rel.}}^2 \rangle [3 \text{ MeV}]^5}}$$

Co-SIMP



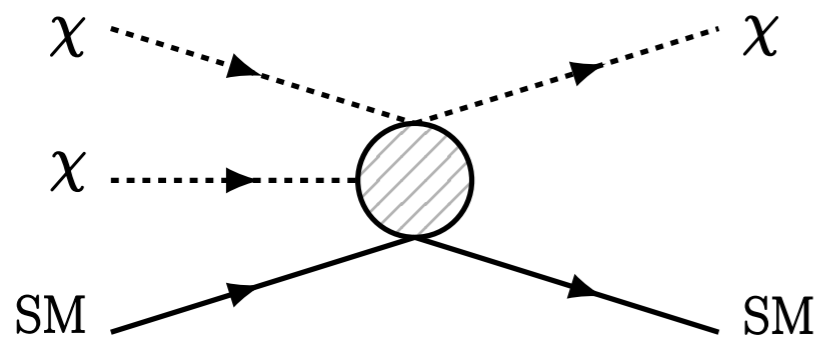
Smirnov, Beacom

$$\Gamma_{\text{DM}} = \langle \sigma_{3 \rightarrow 2} v_{\text{rel.}}^2 \rangle n_{\text{DM}} n_{\text{SM}} > H(T)$$

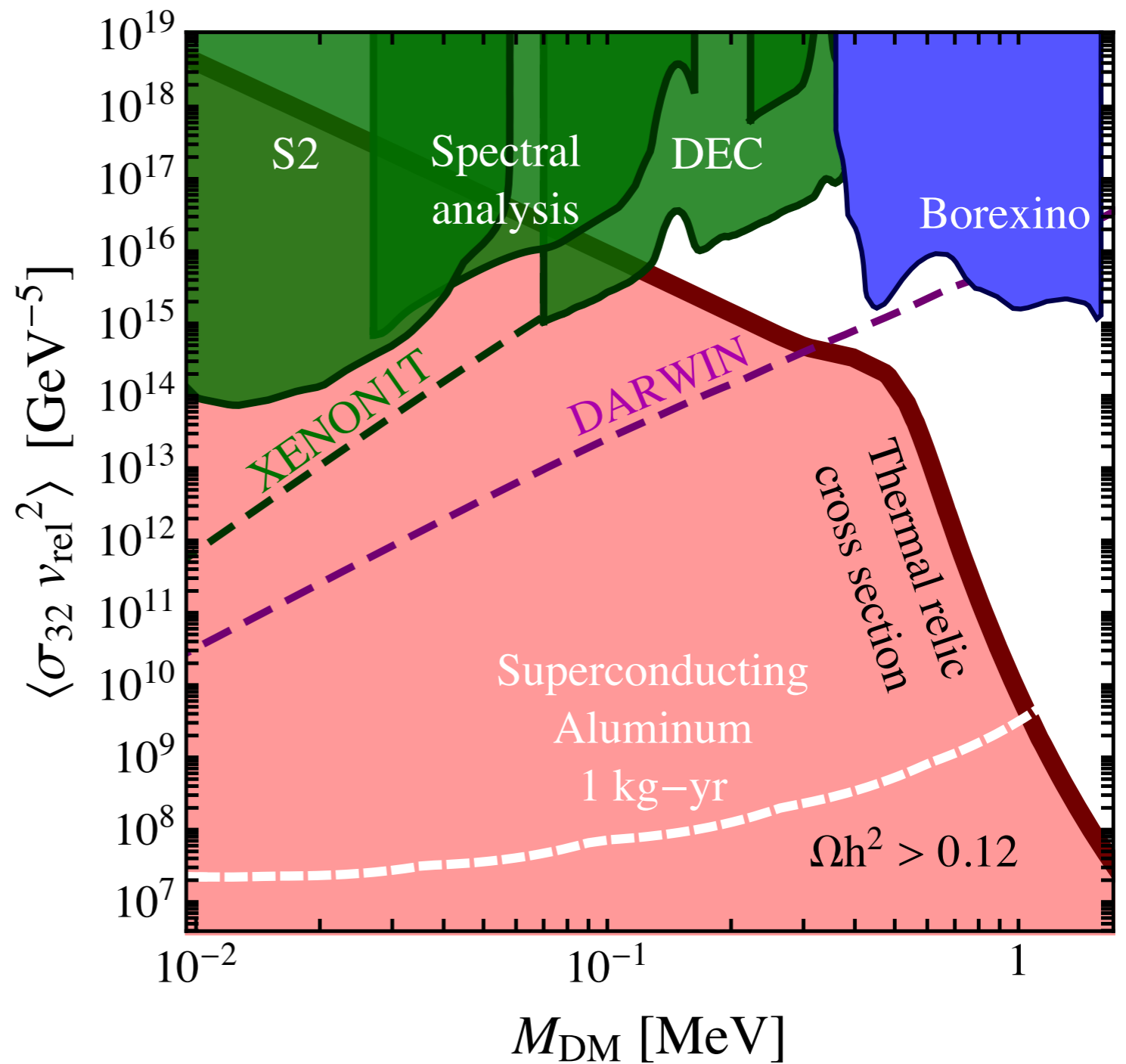
$$\Omega_{\text{DM}} h^2 \approx \left(\frac{m_{\text{DM}}}{\text{GeV}} \right)^3 \frac{0.12}{\langle \sigma_{3 \rightarrow 2} v_{\text{rel.}}^2 \rangle [3 \text{ GeV}]^5}$$

Electrophilic Co-SIMP

Co-SIMP Parameter Space



$$M_\chi < 2M_{\text{el.}}$$

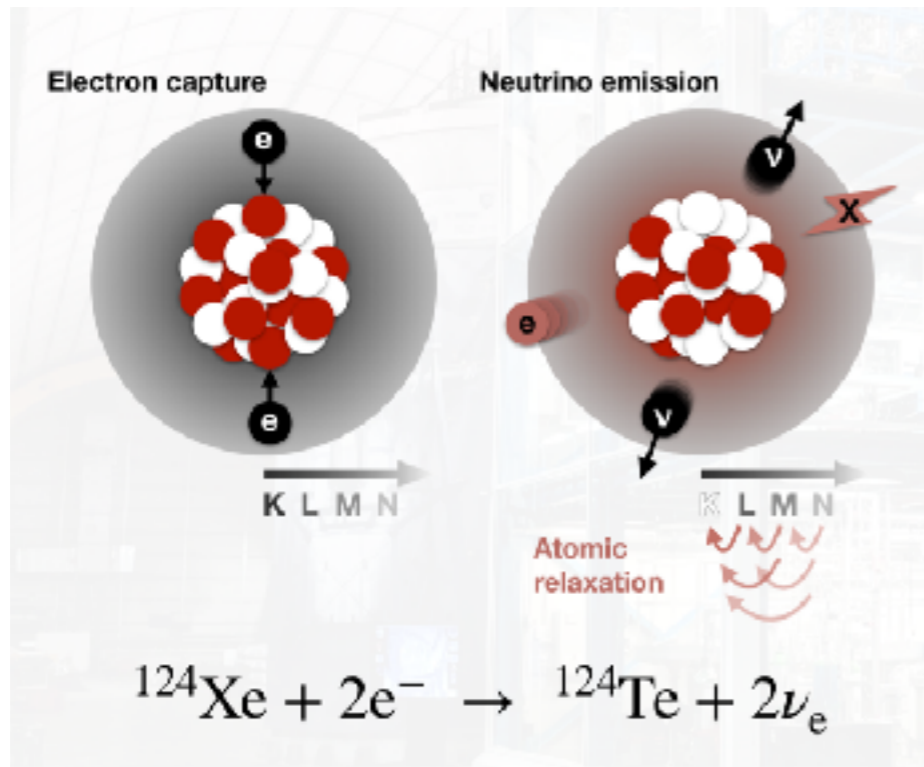


arXiv: 2002.04038; J. Smirnov, J. Beacom

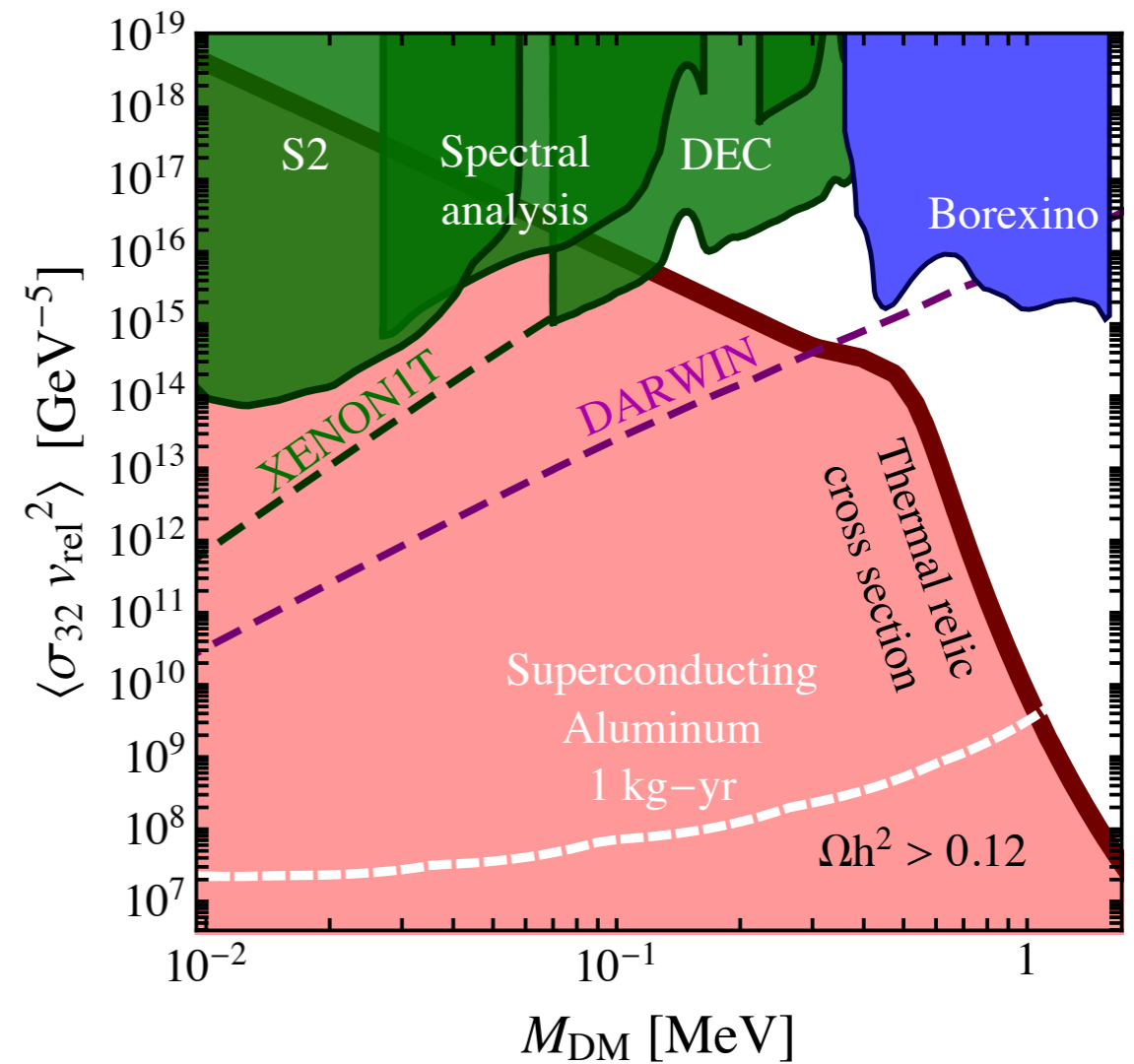
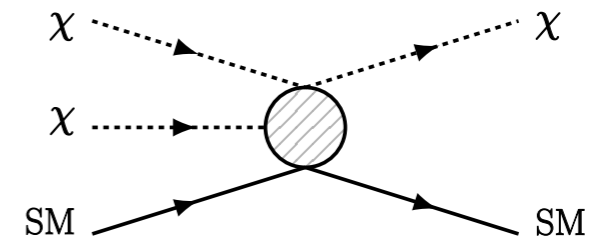
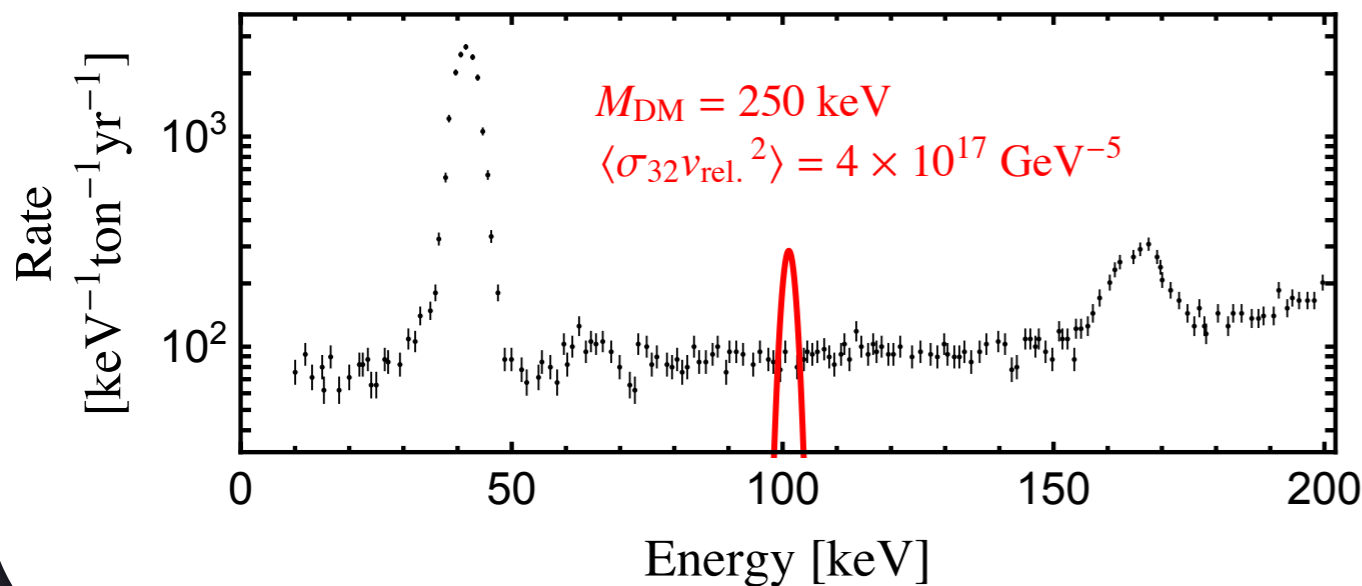
Detection



Double Electron Capture



C. Wittweg, XENON1T Collaboration



arXiv: 2002.04038; J. Smirnov, J. Beacom

We need lower Energy Thresholds

We need lower Energy Thresholds



Superfluid He

We need lower Energy Thresholds



Superfluid He



Materials with Lattice Defects

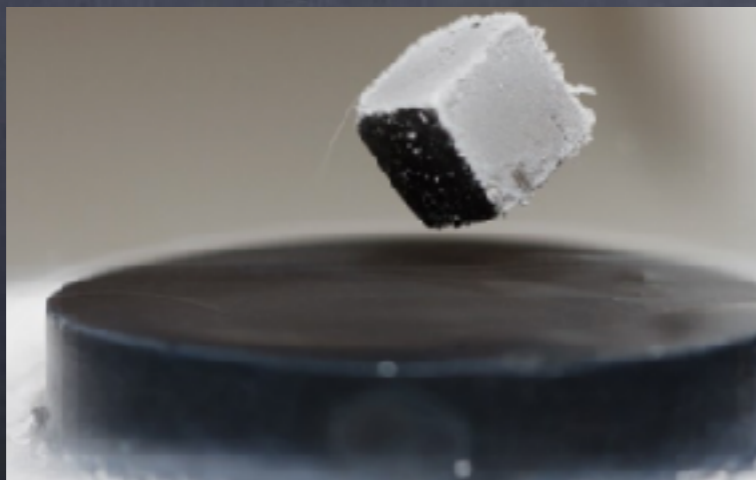
We need lower Energy Thresholds



Superfluid He



Materials with Lattice Defects



Superconductors

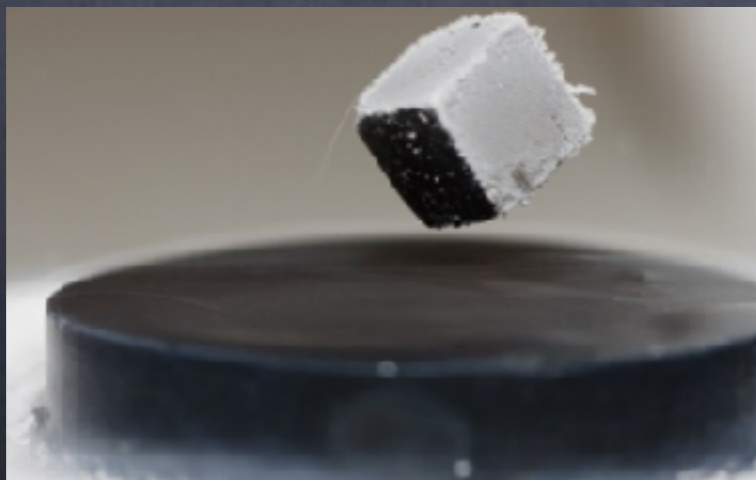
We need lower Energy Thresholds



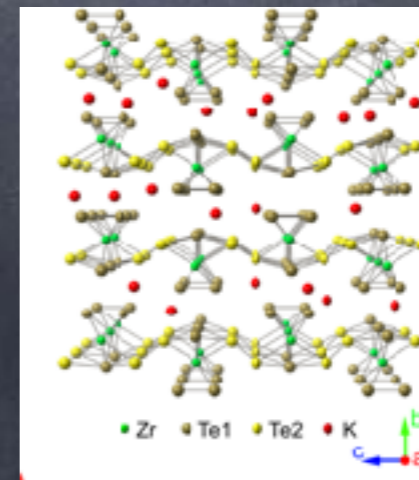
Superfluid He



Materials with Lattice Defects



Superconductors



3D Dirac Materials

We need lower Energy Thresholds

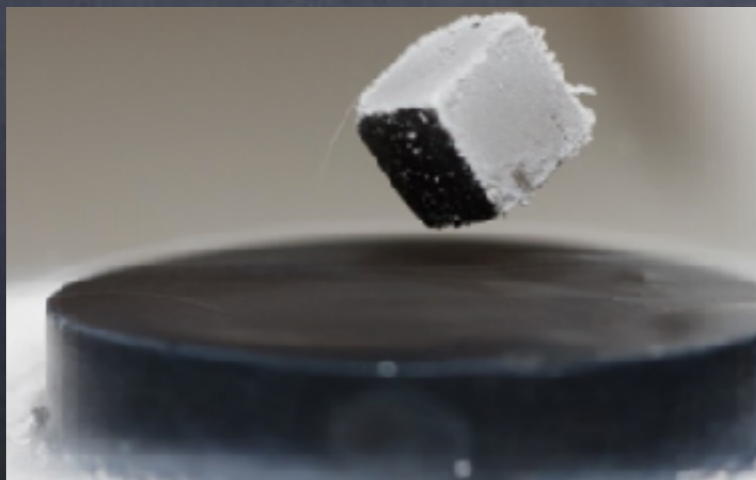


Superfluid He

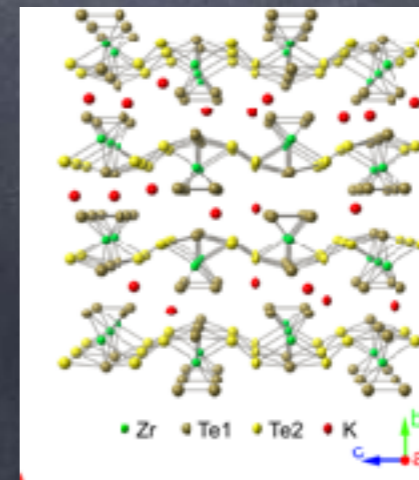


Materials with Lattice Defects

Small ΔE

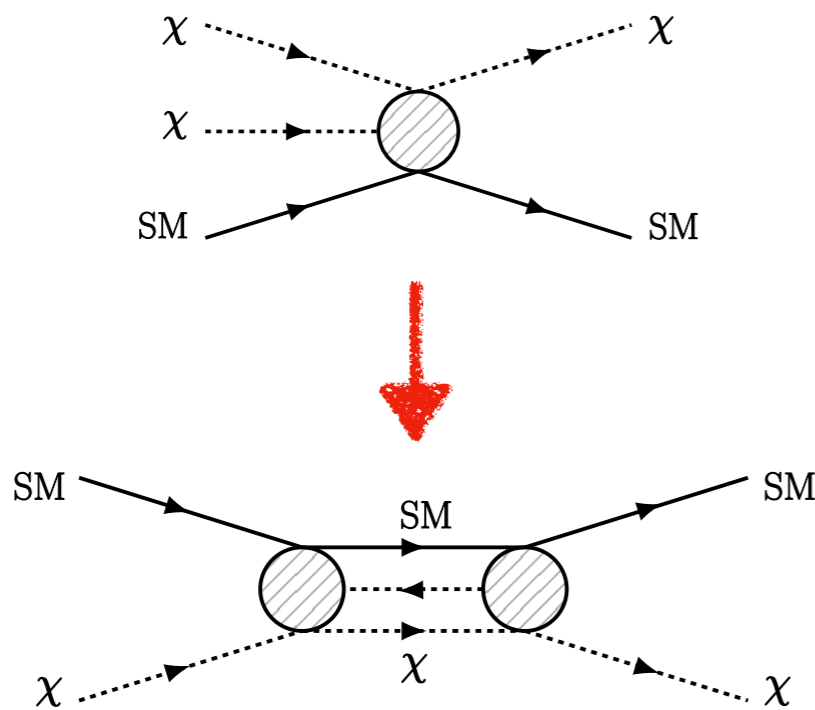


Superconductors



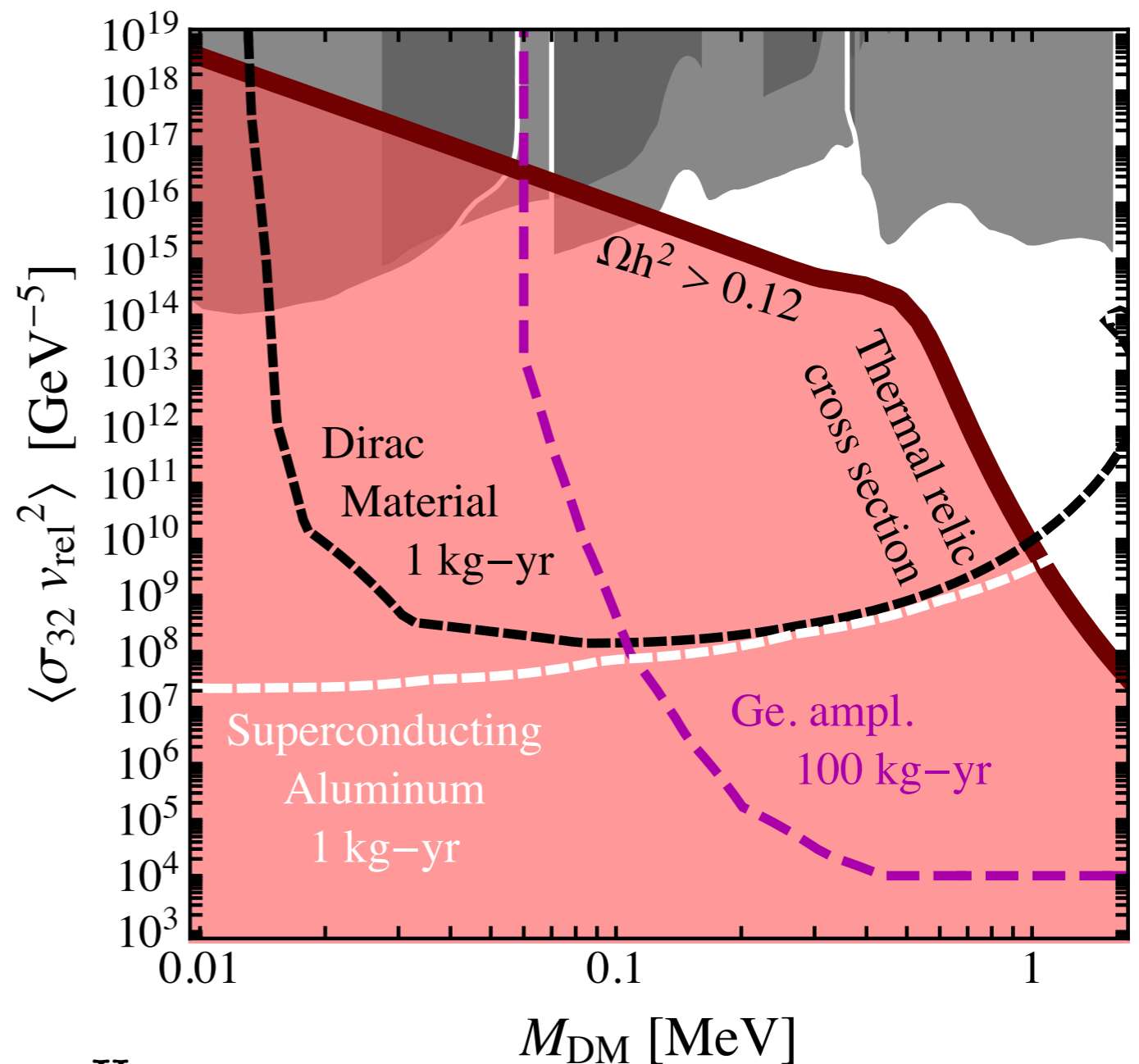
3D Dirac Materials

Scattering off Electrons



$$E_R \approx v_{\text{rel}}^2 \frac{\mu^2}{M_{\text{SM}}}$$

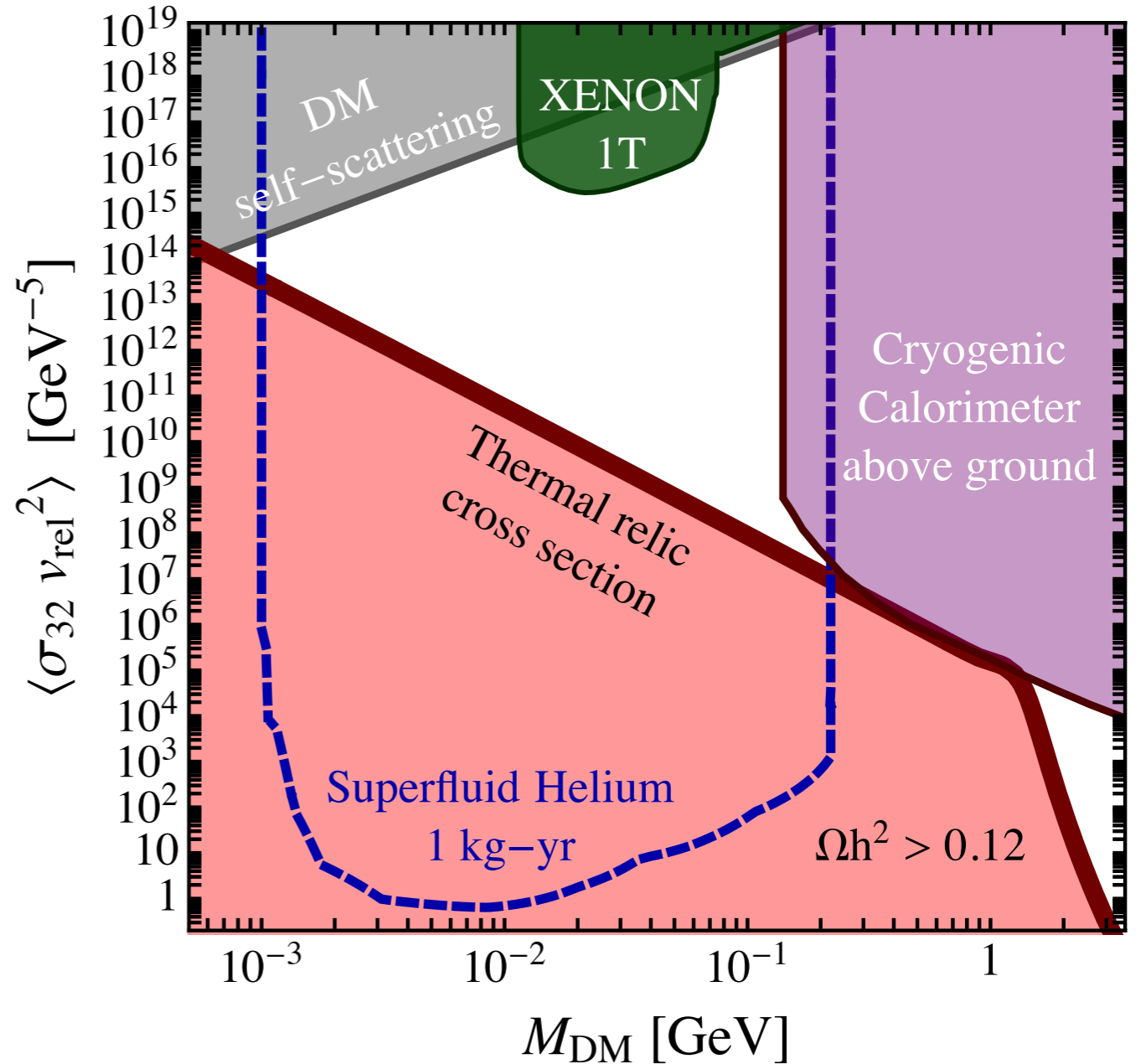
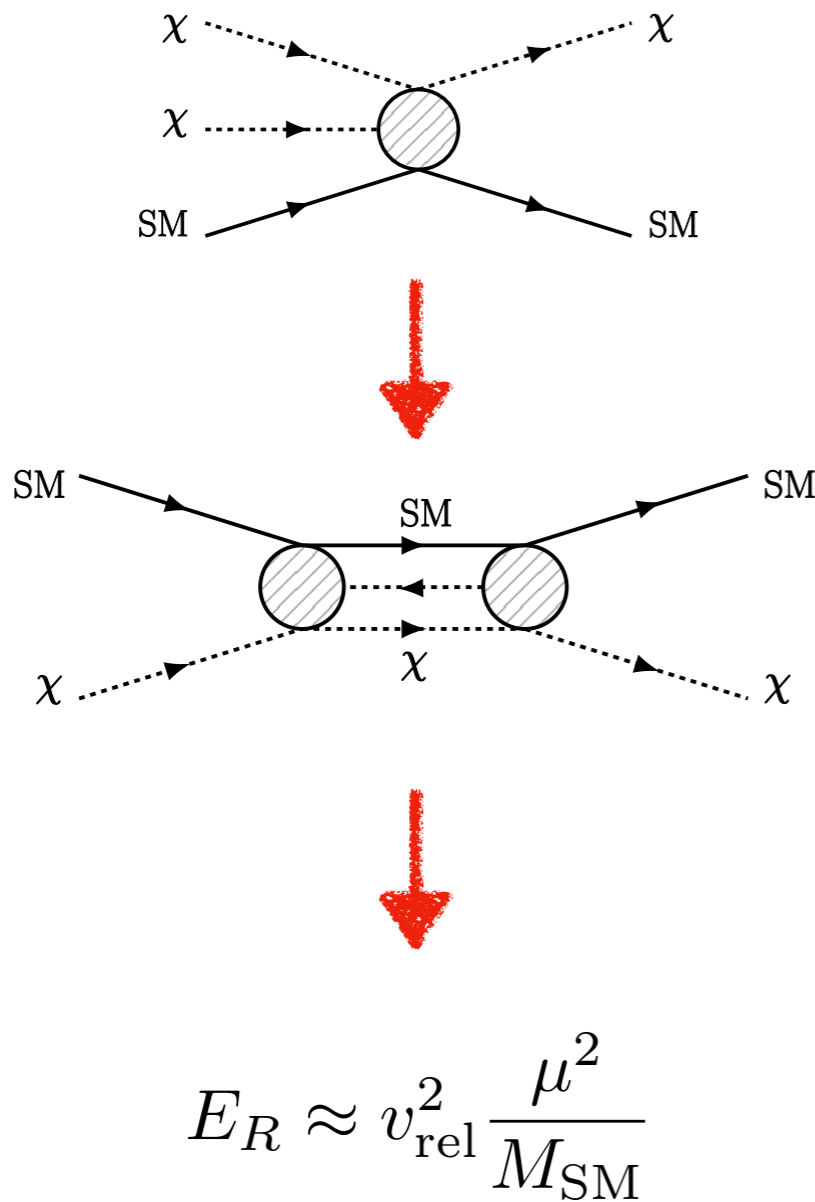
Detectors with lower energy thresholds \approx eV



arXiv: 2002.04038; J. Smirnov, J. Beacom

Nucleophilic Co-SIMP

Scattering off Nucleons



Detectors with lower energy thresholds \approx eV

arXiv: 2002.04038; J. Smirnov, J. Beacom

What next?

What next?

A. Estimates of Supernova Cooling, Co-SIMPs can be trapped. Validate in UV complete model.

What next?

- A. Estimates of Supernova Cooling, Co-SIMPs can be trapped. Validate in UV complete model.
- B. Similar Collider Signatures can only be studied in UV complete models

What next?

- A. Estimates of Supernova Cooling, Co-SIMPs can be trapped. Validate in UV complete model.
- B. Similar Collider Signatures can only be studied in UV complete models
- C. What are possible UV completions? -> Currently working on it with: **Bei Zhou** and **John Beacom** from CCAPP at OSU.

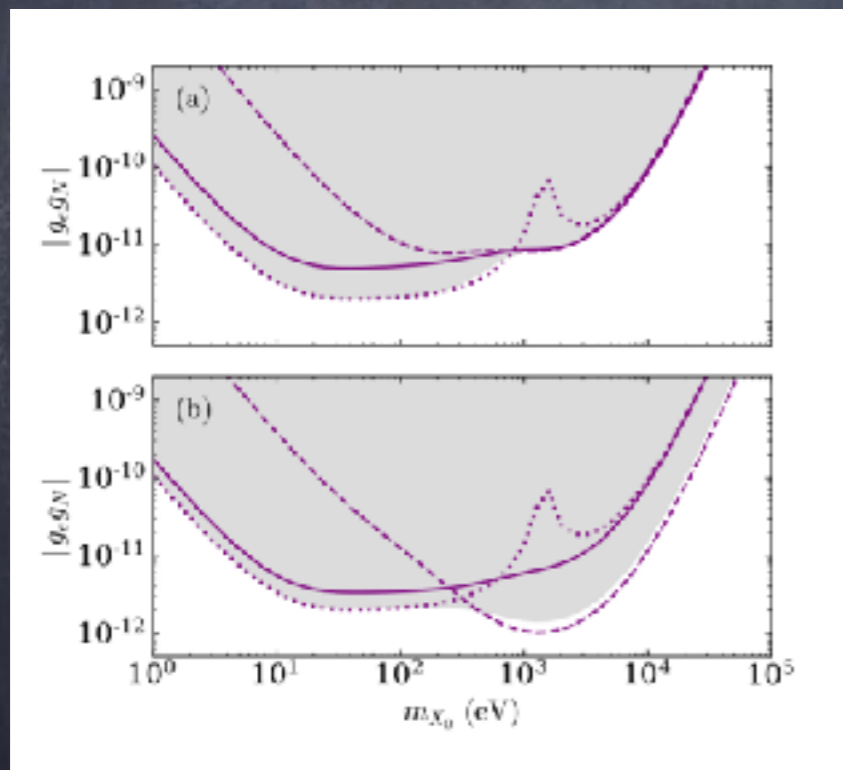
Thanks!

Backup

New Forces

Long Range: light Z'

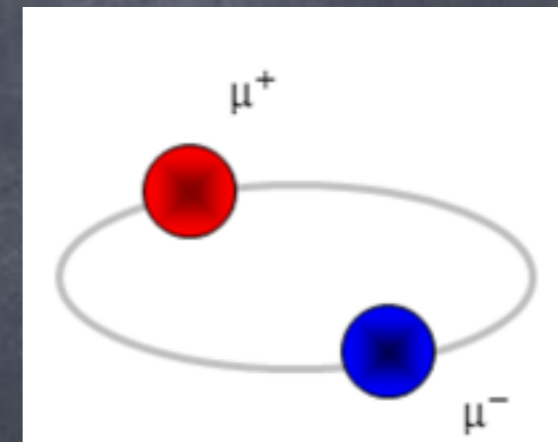
Rydberg Atoms ($n \gg 1$):



arXiv: Jones, Potviliege, Spannowsky

Short Range: loops, heavy Z'

Need compactness:
True Muonium



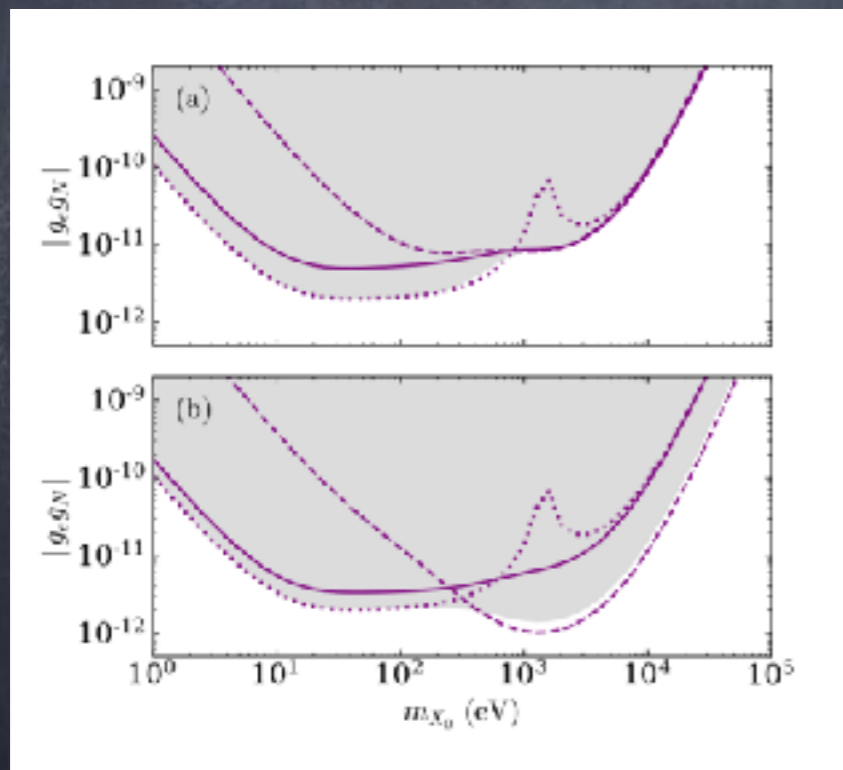
arXiv: Smirnov, Beacom
JCAP 1404 (2014) 022:

Kopp, Michaels, JS

New Forces

Long Range: light Z'

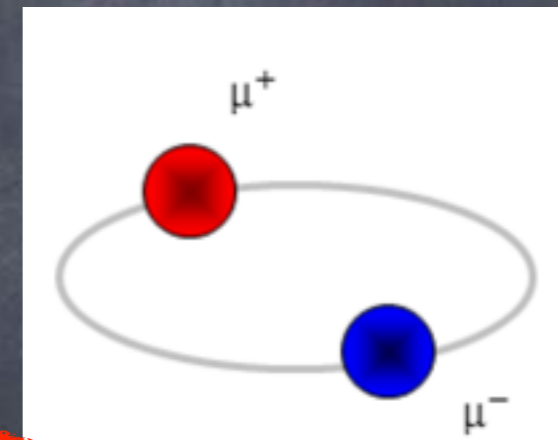
Rydberg Atoms ($n \gg 1$):



arXiv: Jones, Potviliege, Spannowsky

Short Range: loops, heavy Z'

Need compactness:
True Muonium



arXiv: Smirnov, Beacom
JCAP 1404 (2014) 022:

Kopp, Michaels, JS

Positronium

$$\Delta E_{\text{hfs}} \approx -\frac{\alpha_{\text{EM}}^3 m_\ell^3}{8\pi} \frac{1}{128\pi^4 \Lambda^2}$$

JCAP 1404 (2014) 022:

Kopp, Michaels, JS

$$\Delta E_{\text{hfs}}^{e^+e^-} \approx -43 \text{ Hz} \times \left(\frac{\text{GeV}}{\Lambda}\right)^2$$

$$\Delta E/E \approx \mathcal{O}(10^{-10})$$

True Muonium

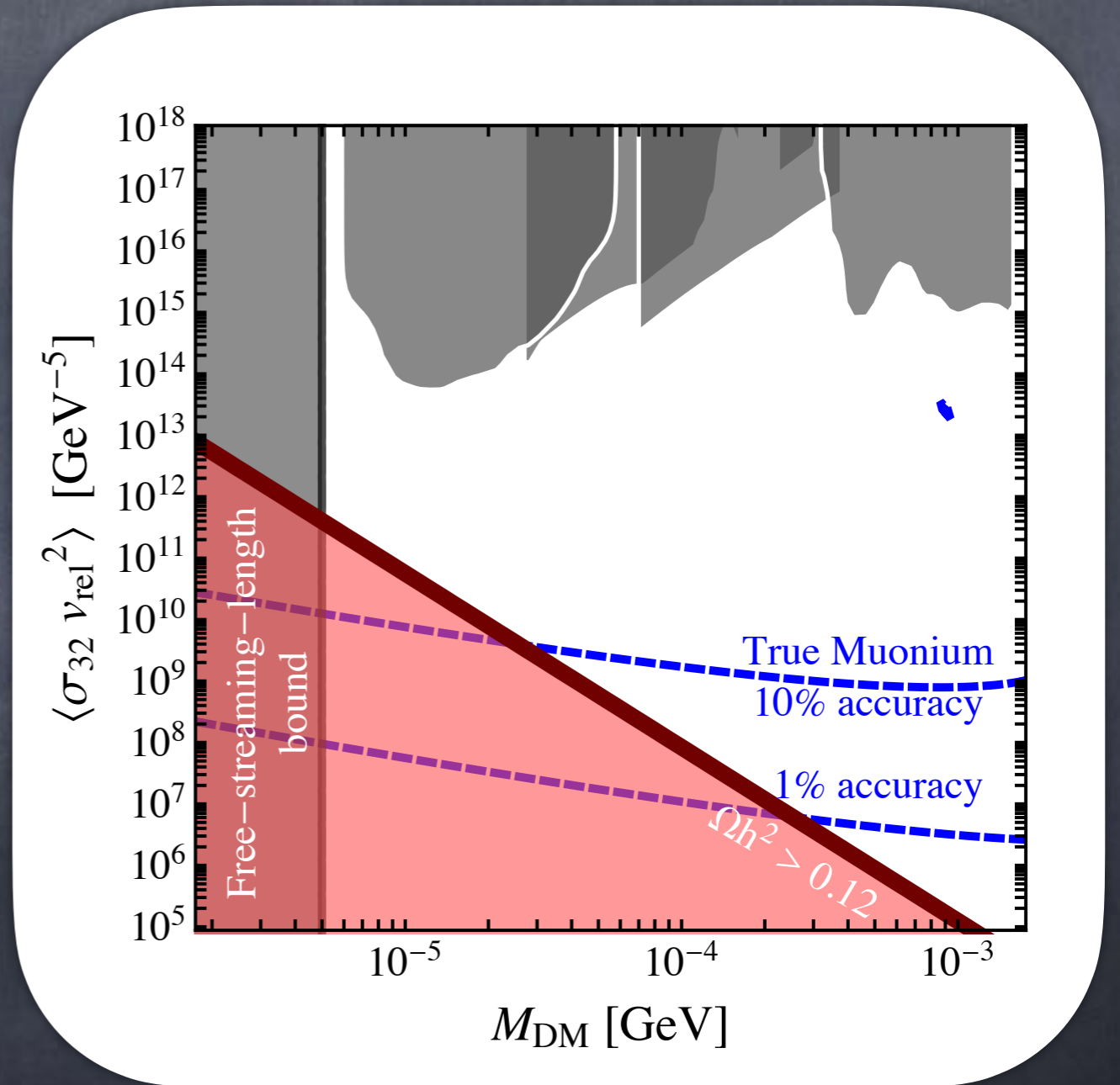
$$\Delta E_{\text{hfs}} \approx -\frac{\alpha_{\text{EM}}^3 m_\ell^3}{8\pi} \frac{1}{128\pi^4 \Lambda^2}$$

JCAP 1404 (2014) 022:
Kopp, Michaels, JS

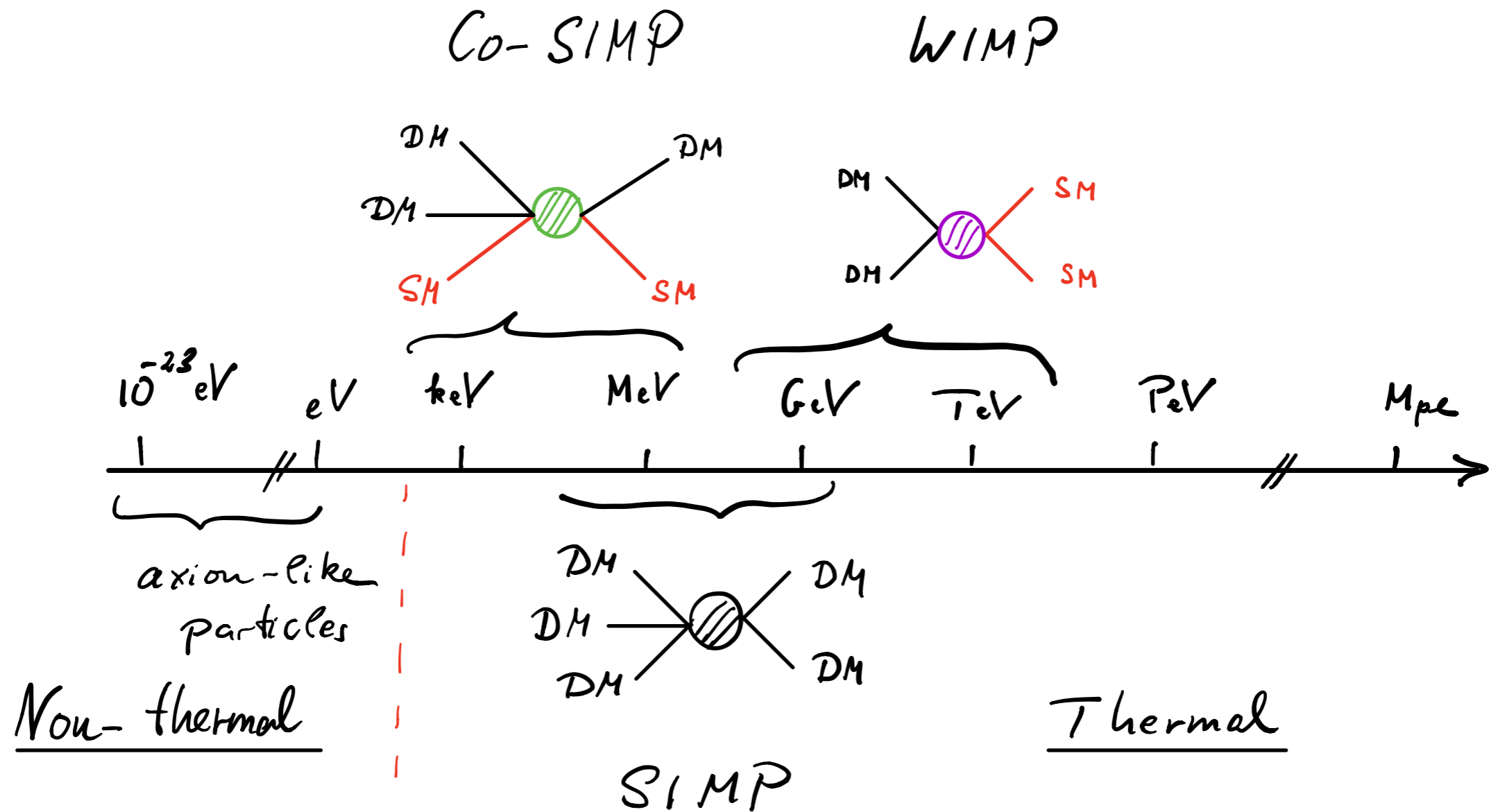
$$\Delta E_{\text{hfs}}^{e^+e^-} \approx -43 \text{ Hz} \times \left(\frac{\text{GeV}}{\Lambda}\right)^2$$

$$\Delta E/E \approx \mathcal{O}(10^{-10})$$

$$\Delta E_{\text{hfs}}^{\mu^+\mu^-} \approx -380 \text{ MHz} \times \left(\frac{\text{GeV}}{\Lambda}\right)^2$$



Dark Particle Production



Example: Dirac Materials

$$H_{\ell} = \begin{pmatrix} 0 & v_F \ell \cdot \boldsymbol{\sigma} - i\Delta \\ v_F \ell \cdot \boldsymbol{\sigma} + i\Delta & 0 \end{pmatrix}, \quad E_{\ell}^{\pm} = \pm \sqrt{v_F^2 \ell^2 + \Delta^2}.$$

