Dark Matter Investigation with DAMA set-ups

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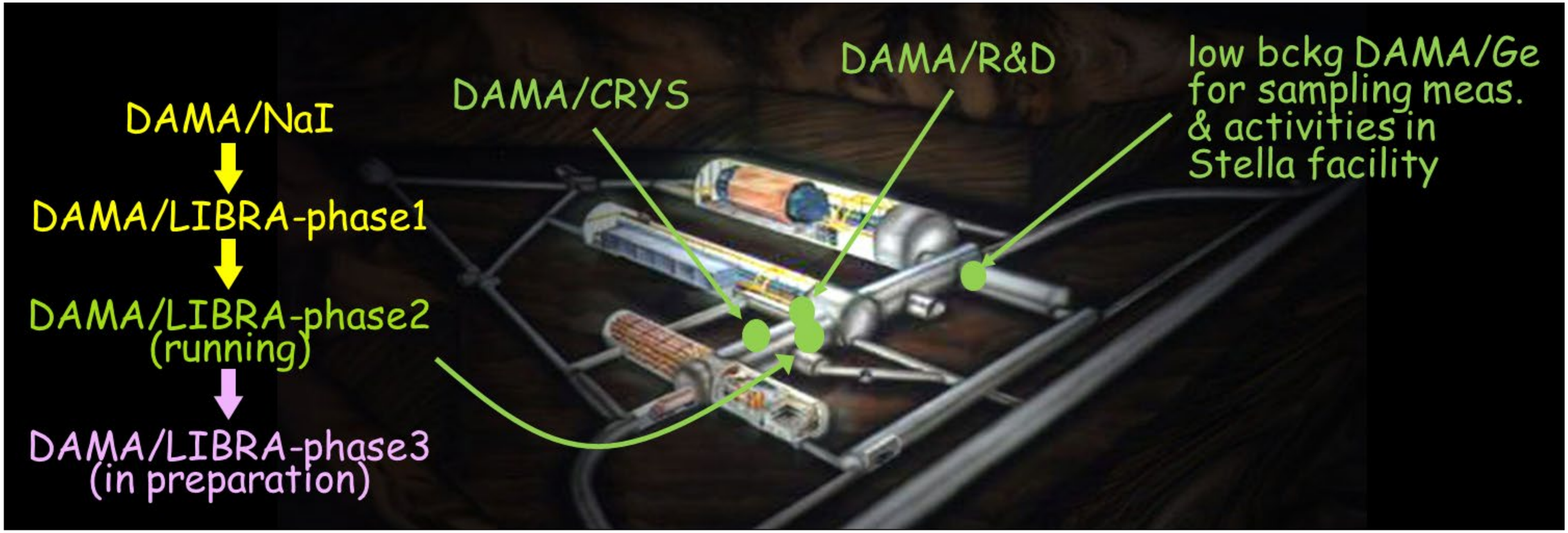
on behalf of DAMA collaboration
University of Roma "Tor Vergata" and INFN

DAMA set-ups

an observatory for rare processes @ LNGS



web site: http://people.roma2.infn.it/dama

Collaboration:

Roma Tor Vergata, Roma La Sapienza, LNGS, IHEP/Beijing

+ by-products and small scale expts.: INR-Kiev + collaborators from other institutions

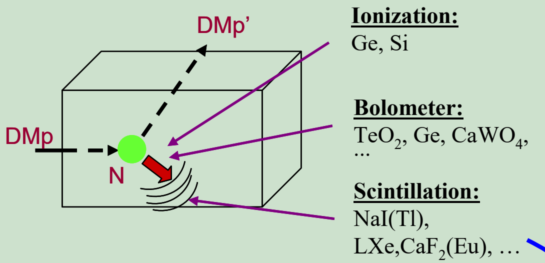
+ neutron meas.: ENEA-Frascati, ENEA-Casaccia

+ in some studies on ββ decays (DST-MAE and Inter-Universities project): IIT Kharagpur and Ropar, India

Some direct detection processes:

- Scatterings on nuclei

  → detection of nuclear recoil energy



- Inelastic Dark Matter: **W + N → W* + N**

  → W has 2 mass states χ+ , χ- with δ mass splitting

  → Kinematical constraint for the inelastic scattering of χ- on a nucleus

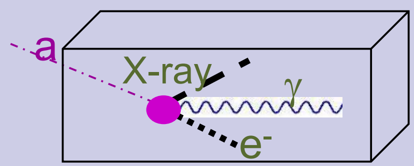
$$\frac{1}{2}\mu v^2 \geq \delta \Leftrightarrow v \geq v_{thr} = \sqrt{\frac{2\delta}{\mu}}$$

- Excitation of bound electrons in scatterings on nuclei

  → detection of recoil nuclei + e.m. radiation

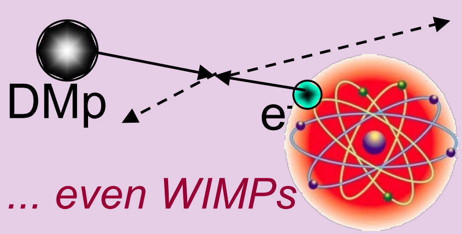
- Conversion of particle into e.m. radiation

  → detection of γ, X-rays, e^-



- Interaction only on atomic electrons

  → detection of e.m. radiation

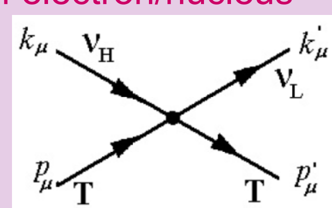


... even WIMPs

- Interaction of light DMp (LDM) on e^- or nucleus with production of a lighter particle

  → detection of electron/nucleus recoil energy

e.g. sterile ν



e.g. signals from these candidates are **completely lost** in experiments based on "rejection procedures" of the e.m. component of their rate

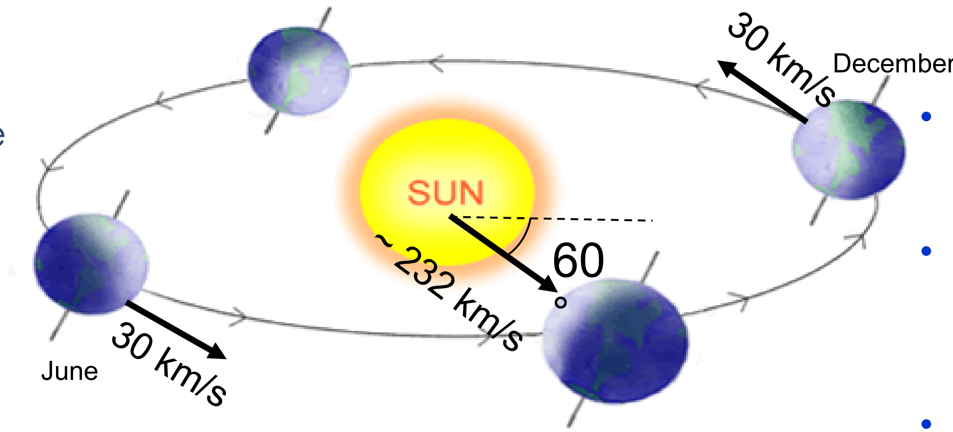
The annual modulation: a model independent signature for the investigation of DM particles component in the galactic halo

With the present technology, the annual modulation is the main model independent signature for the DM signal. Although the modulation effect is expected to be relatively small, a suitable large-mass, low-radioactive set-up with an efficient control of the running conditions can point out its presence.

Drukier, Freese, Spergel PRD86; Freese et al. PRD88

Requirements:

1) Modulated rate according cosine
2) In low energy range
3) With a proper period (1 year)
4) With proper phase (about 2 June)
5) Just for single hit events in a multi-detector set-up
6) With modulation amplitude in the region of maximal sensitivity must be <7% for usually adopted halo distributions, but it can be larger in case of some possible scenarios



- v_{sun} ~ 232 km/s (Sun vel in the halo)
- v_{orb} = 30 km/s (Earth vel around the Sun)
- $\gamma = \pi/3$, $\omega = 2\pi/T$, T = 1 year
- t_0 = 2nd June (when $v_\oplus$ is maximum)

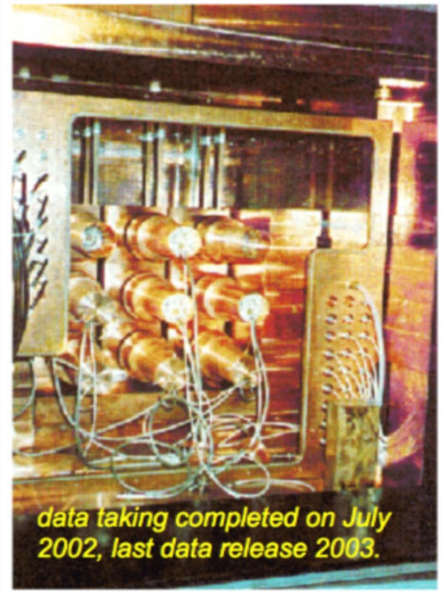
$$v_\oplus(t) = v_{sun} + v_{orb} \cos\gamma \cos[\omega(t-t_0)]$$

$$S_k[\eta(t)] = \int_{\Delta E_k} \frac{dR}{dE_R} dE_R \cong S_{0,k} + S_{m,k} \cos[\omega(t-t_0)]$$

the DM annual modulation signature has a different origin and peculiarities (e.g. the phase) than those effects correlated with the seasons

To mimic this signature, spurious effects and side reactions must not only be able to account for the whole observed modulation amplitude, but also to satisfy contemporaneously all the requirements

The pioneer DAMA/NaI: ≈100 kg highly radiopure NaI(Tl)



Performances:

N.Cim.A112(1999)545-575, EPJC18(2000)283, Riv.N.Cim.26 n. 1(2003)1-73, IJMPD13(2004)2127

Results on rare processes:

- Possible Pauli exclusion principle violation — PLB408(1997)439
- CNC processes — PRC60(1999)065501
- Electron stability and non-paulian transitions in Iodine atoms (by L-shell) — PLB460(1999)235
- Search for solar axions — PLB515(2001)6
- Exotic Matter search — EPJdirect C14(2002)1
- Search for superdense nuclear matter — EPJA23(2005)7
- Search for heavy clusters decays — EPJA24(2005)51

Results on DM particles:

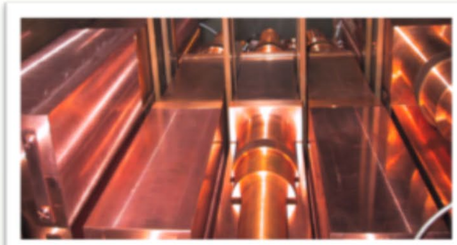
- PSD — PLB389(1996)757
- Investigation on diurnal effect — N.Cim.A112(1999)1541
- Exotic Dark Matter search — PRL83(1999)4918
- **Annual Modulation Signature** — PLB424(1998)195, PLB450(1999)448, PRD61(1999)023512, PLB480(2000)23, EPJC18(2000)283, PLB509(2001)197, EPJC23(2002)61, PRD66(2002)043503, Riv.N.Cim.26 n.1 (2003)1, IJMPD13(2004)2127, IJMPA21(2006)1445, EPJC47(2006)263, IJMPA22(2007)3155, EPJC53(2008)205, PRD77(2008)023506, MPLA23(2008)2125

Model independent evidence of a particle DM component in the galactic halo at 6.3σ C.L.

total exposure (7 annual cycles) 0.29 ton×yr

The DAMA/LIBRA set-up ~250 kg NaI(Tl) (Large sodium Iodide Bulk for RAre processes)

As a result of a 2nd generation R&D for more radiopure NaI(Tl) by exploiting new chemical/physical radiopurification techniques
(all operations involving - including photos - in HP Nitrogen atmosphere)



Residual contaminations in the new DAMA/LIBRA NaI(Tl) detectors: ^{232}Th, ^{238}U and ^{40}K at level of 10^{-12} g/g

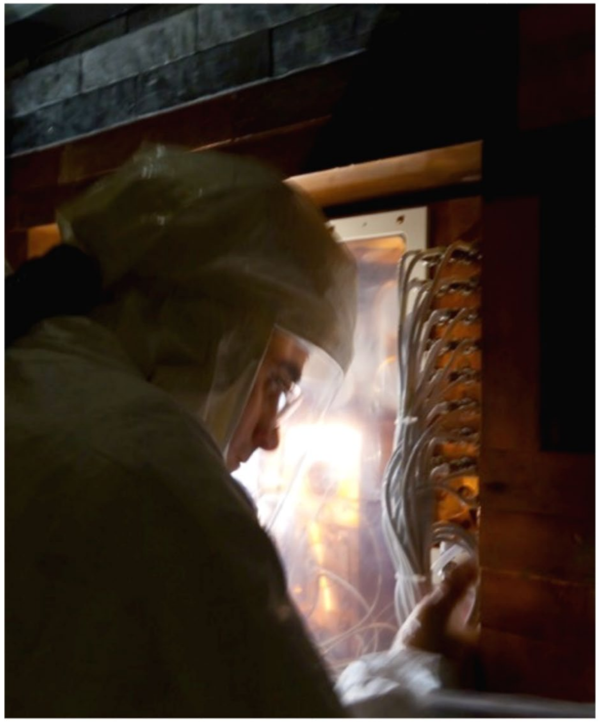
- Radiopurity, performances, procedures, etc.: NIMA592(2008)297, JINST 7 (2012) 03009
- Results on DM particles,
  - Annual Modulation Signature: EPJC56(2008)333, EPJC67(2010)39, EPJC73(2013)2648.
  - Related results: PRD84(2011)055014, EPJC72(2012)2064, IJMPA28(2013)1330022, EPJC74(2014)2827, EPJC74(2014)3196, EPJC75(2015)239, EPJC75(2015)400, IJMPA31(2016) dedicated issue, EPJC77(2017)83
- Results on rare processes:
  - PEPv: EPJC62(2009)327, arXiv1712.08082;
  - CNC: EPJC72(2012)1920;
  - IPP in ^{241}Am: EPJA49(2013)64

DAMA/LIBRA–phase1 (7 annual cycles, 1.04 ton×yr) confirmed the model-independent evidence of DM: reaching 9.3σ C.L.

DAMA/LIBRA–phase2

Upgrade on Nov/Dec 2010: all PMTs replaced with new ones of higher Q.E.

JINST 7(2012)03009
Universe 4 (2018) 116
NPAE 19 (2018) 307
Bled 19 (2018) 27
NPAE 20(4)(2019)317
N.Cim. C 43 (2020) 23
PPNP 114(2020)103810



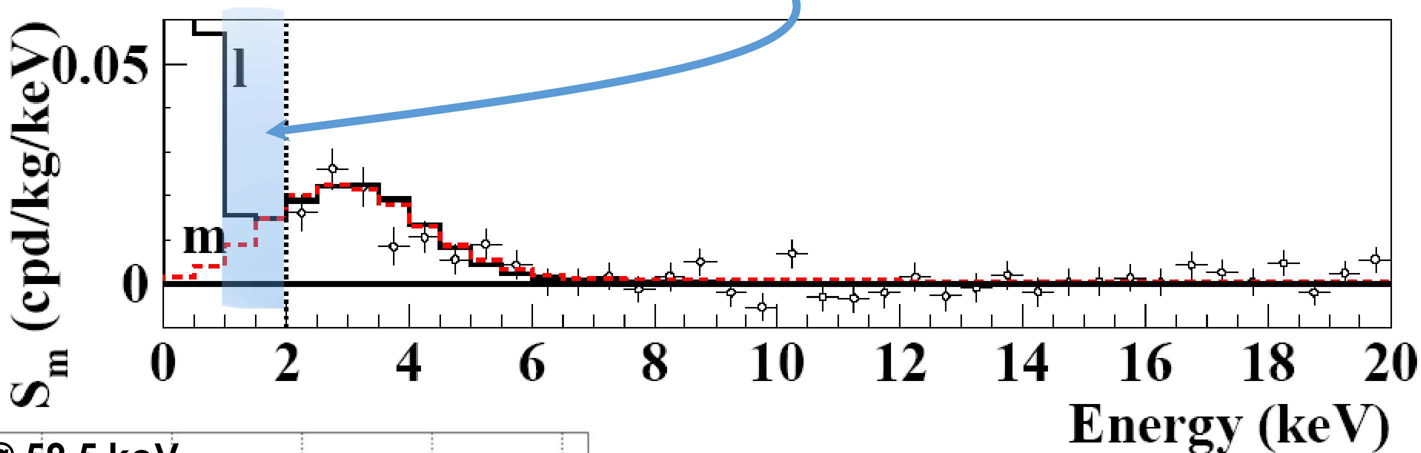
Q.E. of the new PMTs:
33 – 39% @ 420 nm
36 – 44% @ peak

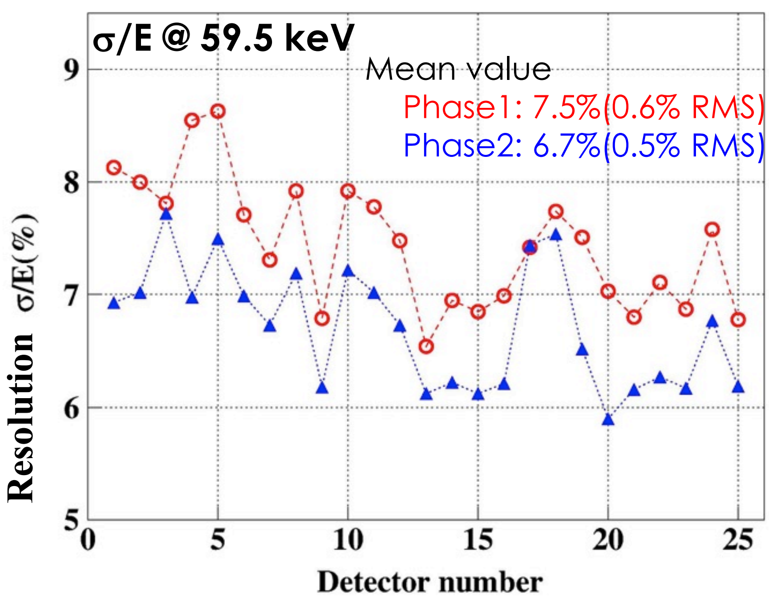
DAMA/LIBRA–phase2

JINST 7(2012)03009
Universe 4 (2018) 116
NPAE 19 (2018) 207
Bled W. in Phys.19 (2018) 27
PPNP 114(2020)103810

Lowering software energy threshold below 2 keV:

- to study the nature of the particles and features of astrophysical, nuclear and particle physics aspects, and to investigate 2nd order effects
- special data taking for *other rare processes*





PMTs contaminations:

	^{226}Ra (Bq/kg)	^{235}U (mBq/kg)	^{228}Ra (Bq/kg)	^{228}Th (mBq/kg)	^{40}K (Bq/kg)
Mean Contamination	0.43	47	0.12	83	0.54
Standard Deviation	0.06	10	0.02	17	0.16

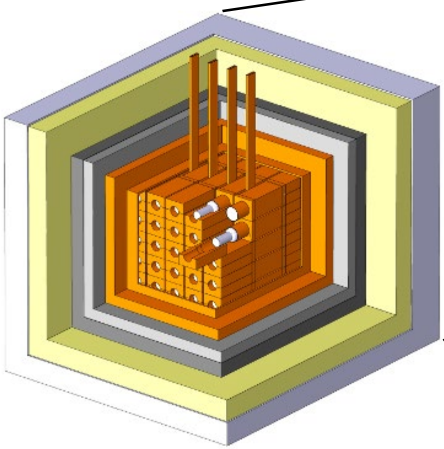
The light responses:

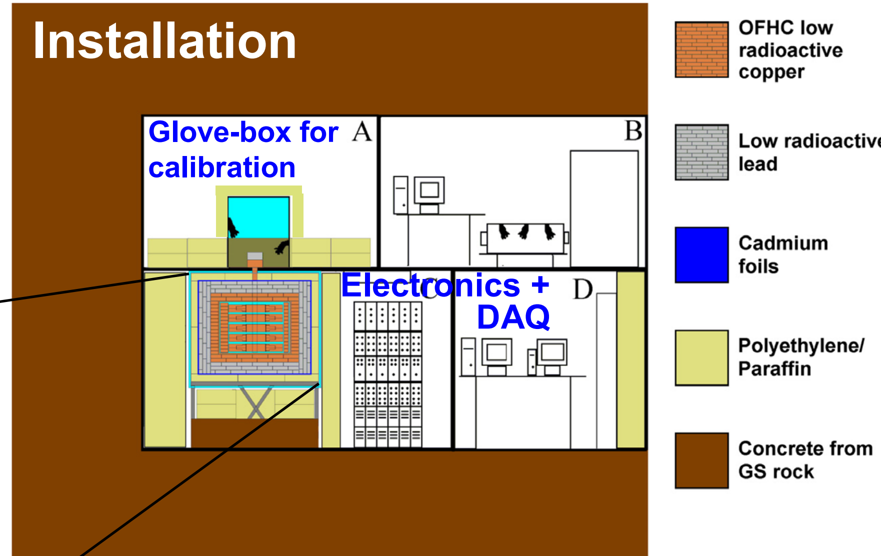
DAMA/LIBRA-phase1: 5.5 – 7.5 ph.e./keV
DAMA/LIBRA-phase2: 6-10 ph.e./keV

The DAMA/LIBRA–phase2 set-up

NIMA592(2008)297, JINST 7(2012)03009, IJMPA31(2017)issue31

- 25 x 9.7 kg NaI(Tl) in a 5x5 matrix
- two Suprasil-B light guides directly coupled to each bare crystal
- two new high Q.E. PMTs for each crystal working in coincidence at the single ph. el. threshold
- **6-10 phe/keV; 1 keV software energy threshold**

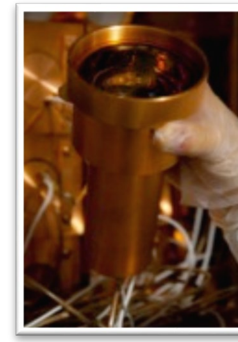




- Multiton-multicomponent passive shield (>10 cm OFHC Cu, 15 cm boliden Pb + Cd foils, 10/40 cm polyethylene/paraffin, ~1 m concrete, mostly outside the installation)
- Three-level system to exclude Radon from the detectors
- Calibrations in the same running conditions as prod runs
- Never neutron source in DAMA installations
- Installation in air conditioning + huge heat capacity of shield
- Monitoring/alarm system; many parameters acquired with the production data

- Whole setup decoupled from ground
- Fragmented set-up: single-hit events = each detector has all the others as anticoincidence
- Dismounting/Installing protocol in HP N_2
- All the materials selected for low radioactivity
- Pulse shape recorded by Waweform Analyzer Acqiris DC270 (2chs per detector), 1 *Gs/s,* 8 bit, bandwidth 250 MHz both for single-hit and multiple-hit events
- Data collected from low energy up to MeV region, despite the hardware optimization for low energy
- DAQ with optical readout
- New electronic modules

DAMA/LIBRA-phase2 data taking



Second upgrade at end of 2010: all PMTs replaced with new ones of higher Q.E.

JINST 7(2012)03009

Energy resolution @ 60 keV mean value: prev. PMTs 7.5% (0.6% RMS) new HQE PMTs 6.7% (0.5% RMS)

- ✓ Fall 2012: new preamplifiers installed + special trigger modules.
- ✓ Calibrations 6 a.c.: ≈ 1.3×10^8 events from sources
- ✓ Acceptance window eff. 6 a.c.: ≈ 3.4×10^6 events (≈1.4×10^5 events/keV)

Annual Cycles	Period	Mass (kg)	Exposure (kg×day)	$(\alpha-\beta^2)$
I	Dec 23, 2010 – Sept. 9, 2011	commissioning		
II	Nov. 2, 2011 – Sept. 11, 2012	242.5	62917	0.519
III	Oct. 8, 2012 – Sept. 2, 2013	242.5	60586	0.534
IV	Sept. 8, 2013 – Sept. 1, 2014	242.5	73792	0.479
V	Sept. 1, 2014 – Sept. 9, 2015	242.5	71180	0.486
VI	Sept. 10, 2015 – Aug. 24, 2016	242.5	67527	0.522
VII	Sept. 7, 2016 – Sept. 25, 2017	242.5	75135	0.480

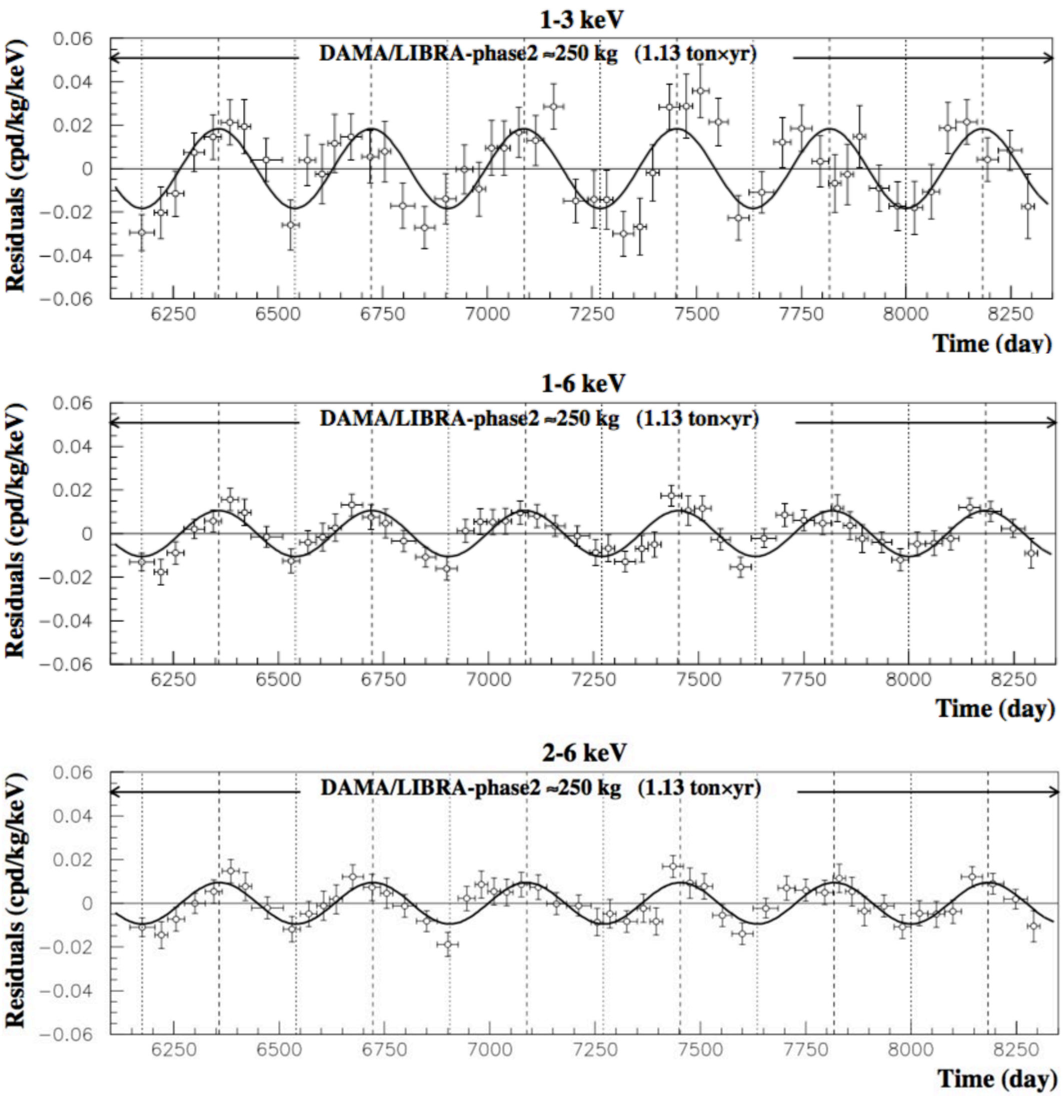
Exposure first data release of DAMA/LIBRA-phase2: **1.13 ton × yr**

Exposure DAMA/NaI+DAMA/LIBRA-phase1+phase2: **2.46 ton × yr**

DM model-independent Annual Modulation Result

Experimental residuals of the single-hit scintillation events rate vs time and energy

DAMA/LIBRA-phase2 (1.13 ton×yr)



Absence of modulation? No

- 1-3 keV: χ^2/dof=127/52 $\Rightarrow$ P(A=0) = 3×10^{-8}
- 1-6 keV: χ^2/dof=150/52 $\Rightarrow$ P(A=0) = 2×10^{-11}
- 2-6 keV: χ^2/dof=116/52 $\Rightarrow$ P(A=0) = 8×10^{-7}

Fit on DAMA/LIBRA-phase2

$A\cos[\omega(t-t_0)]$;
continuous lines: t_0 = 152.5 d, T = 1.00 y

1-3 keV

A=(0.0184±0.0023) cpd/kg/keV

χ^2/dof = 61.3/51 **8.0 σ C.L.**

1-6 keV

A=(0.0105±0.0011) cpd/kg/keV

χ^2/dof = 50.0/51 **9.5 σ C.L.**

2-6 keV

A=(0.0095±0.0011) cpd/kg/keV

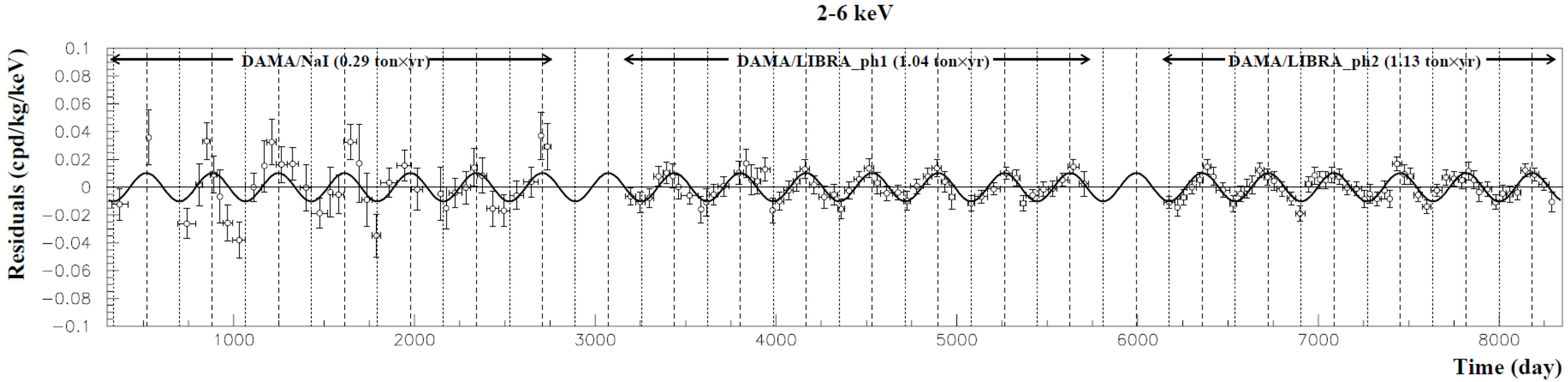
χ^2/dof = 42.5/51 **8.6 σ C.L.**

The data of DAMA/LIBRA-phase2 favor the presence of a modulated behavior with proper features at 9.5σ C.L.

DM model-independent Annual Modulation Result

Experimental residuals of the single-hit scintillation events rate vs time and energy

DAMA/NaI+DAMA/LIBRA-phase1+DAMA/LIBRA-phase2 (2.46 ton × yr)



Absence of modulation? No

- 2-6 keV: χ^2/dof=272.3/142 $\Rightarrow$ P(A=0) =3.0×10^{-10}

Fit on DAMA/NaI+ DAMA/LIBRA-ph1+ DAMA/LIBRA-ph2

$A\cos[\omega(t-t_0)]$;
continuous lines: t_0 = 152.5 d, T = 1.00 y

2-6 keV

A=(0.0102±0.0008) cpd/kg/keV

χ^2/dof = 113.8/138 **12.8 σ C.L.**

The data of DAMA/NaI + DAMA/LIBRA-phase1 +DAMA/LIBRA-phase2 favor the presence of a modulated behavior with proper features at 12.8 σ C.L.

Releasing period (T) and phase (t_0) in the fit

	ΔE	A(cpd/kg/keV)	$T=2\pi/\omega$ (yr)	t_0 (day)	C.L.
DAMA/LIBRA-ph2	(1-3) keV	0.0184±0.0023	1.0000±0.0010	153±7	8.0σ
	(1-6) keV	0.0106±0.0011	0.9993±0.0008	148±6	9.6σ
	(2-6) keV	0.0096±0.0011	0.9989±0.0010	145±7	8.7σ
DAMA/LIBRA-ph1 + DAMA/LIBRA-ph2	(2-6) keV	0.0096±0.0008	0.9987±0.0008	145±5	12.0σ
DAMA/NaI + DAMA/LIBRA-ph1 + DAMA/LIBRA-ph2	(2-6) keV	0.0103±0.0008	0.9987±0.0008	145±5	12.9σ

$A\cos[\omega(t-t_0)]$

DAMA/NaI (0.29 ton x yr)

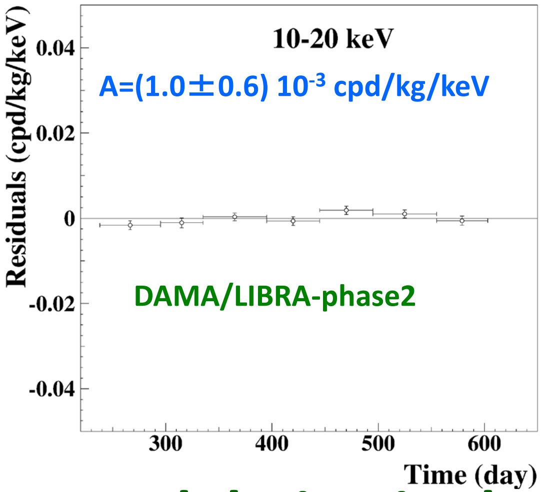
DAMA/LIBRA-ph1 (1.04 ton x yr)

DAMA/LIBRA-ph2 (1.13 ton x yr)

total exposure = 2.46 ton×yr

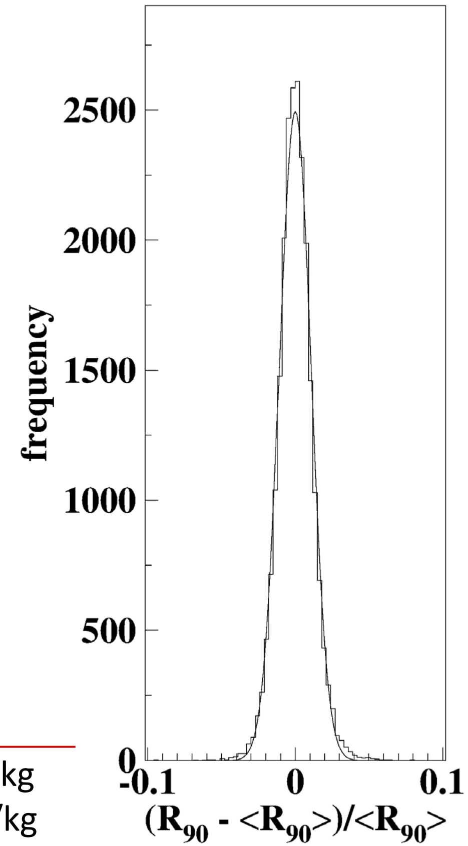
Rate behaviour above 6 keV

• No Modulation above 6 keV



Mod. Ampl. (6-14 keV): cpd/kg/keV
(0.0032 ± 0.0017) DAMA/LIBRA-ph2_2
(0.0016 ± 0.0017) DAMA/LIBRA-ph2_3
(0.0024 ± 0.0015) DAMA/LIBRA-ph2_4
-(0.0004 ± 0.0015) DAMA/LIBRA-ph2_5
(0.0001 ± 0.0015) DAMA/LIBRA-ph2_6
(0.0015 ± 0.0014) DAMA/LIBRA-ph2_7
→ statistically consistent with zero

DAMA/LIBRA-phase2



$\sigma \approx 1\%$, fully accounted by statistical considerations

• No modulation in the whole energy spectrum:

studying integral rate at higher energy, R_{90}

- R_{90} percentage variations with respect to their mean values for single crystal
- Fitting the behaviour with time, adding a term modulated with period and phase as expected for DM particles: **consistent with zero**

Period	Mod. Ampl.
DAMA/LIBRA-ph2_2	(0.12±0.14) cpd/kg
DAMA/LIBRA-ph2_3	-(0.08±0.14) cpd/kg
DAMA/LIBRA-ph2_4	(0.07±0.15) cpd/kg
DAMA/LIBRA-ph2_5	-(0.05±0.14) cpd/kg
DAMA/LIBRA-ph2_6	(0.03±0.13) cpd/kg
DAMA/LIBRA-ph2_7	-(0.09±0.14) cpd/kg

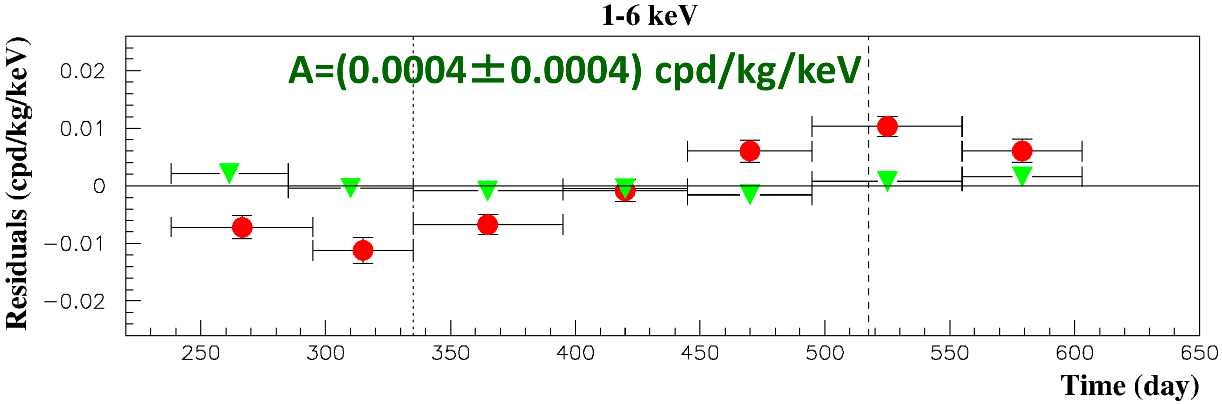
+ if a modulation present in the whole energy spectrum at the level found in the lowest energy region → R_{90} ~ tens cpd/kg → ~ 100 σ far away

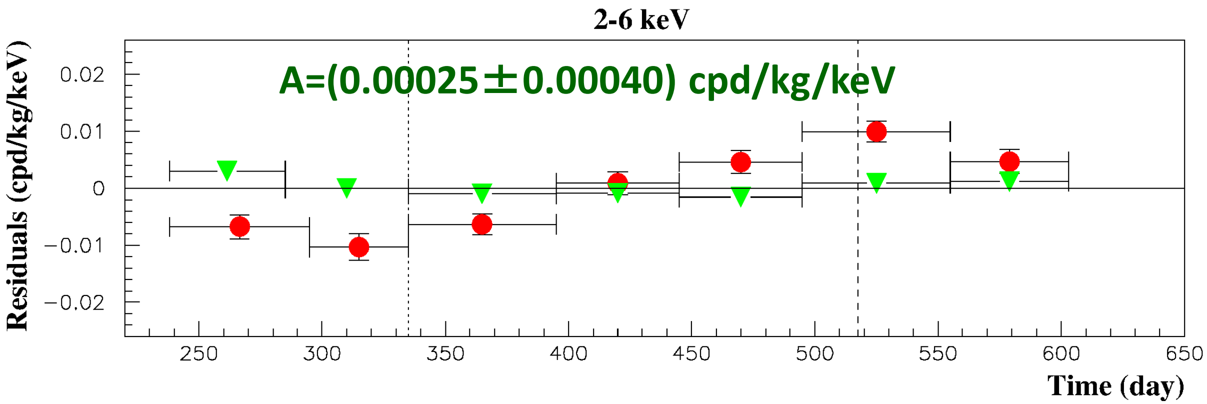
No modulation above 6 keV
This accounts for all sources of background and is consistent with the studies on the various components

DM model-independent Annual Modulation Result

DAMA/LIBRA-phase2 (1.13 ton × yr)

Multiple hits events = Dark Matter particle "switched off"





Single hit residual rate (red)

vs

Multiple hit residual rate (green)

- Clear modulation in the single hit events;
- No modulation in the residual rate of the multiple hit events

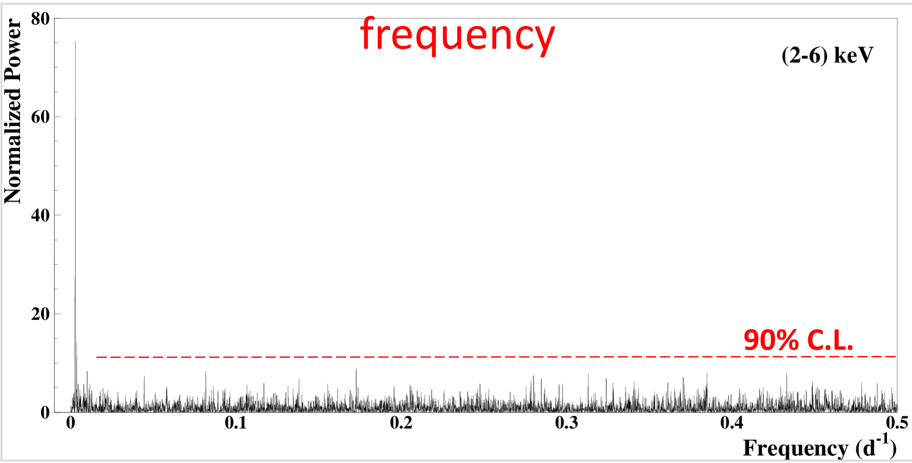
This result furthermore rules out any side effect either from hardware or from software procedures or from background

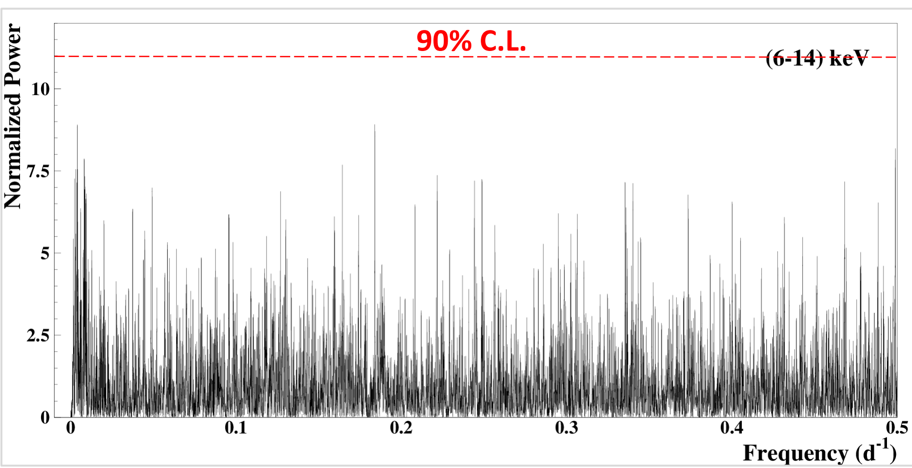
The analysis in frequency

(according to PRD75 (2007) 013010)

To perform the Fourier analysis of the data in a wide region of frequency, the single-hit scintillation events have been grouped in **1 day bins**

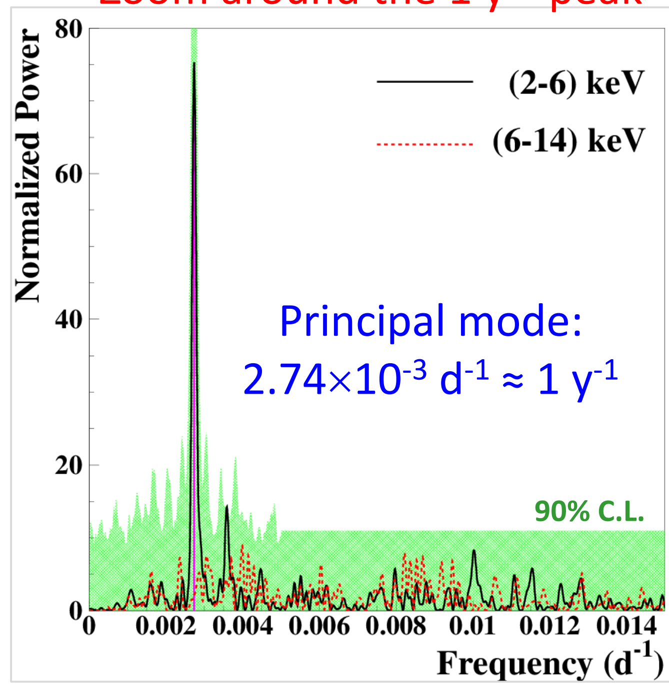
The whole power spectra up to the Nyquist frequency





DAMA/NaI + DAMA/LIBRA-(ph1+ph2) (20 yr)
total exposure: 2.46 ton×yr

Zoom around the 1 y^{-1} peak



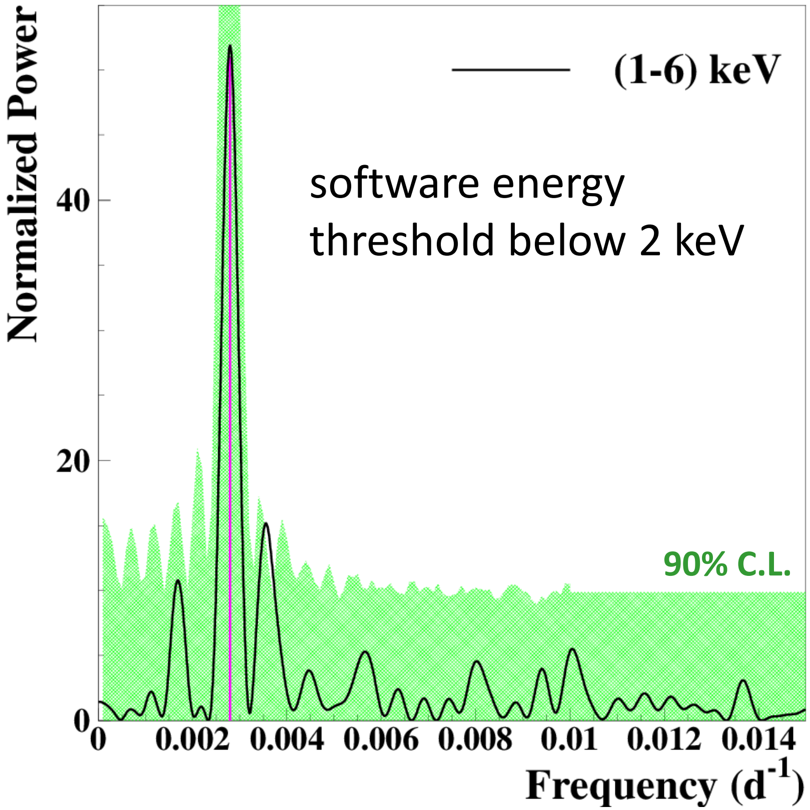
Green area: 90% C.L. region calculated taking into account the signal in (2-6) keV

Clear annual modulation in (2-6) keV + only aliasing peaks far from signal region

The analysis in frequency

(according to PRD75 (2007) 013010)

To perform the Fourier analysis of the data in a wide region of frequency, the single-hit scintillation events have been grouped in 1 day bins



DAMA/LIBRA-phase2 (6 yr)
total exposure: 1.13 ton×yr

Principal mode: 2.79×10^{-3} d^{-1} ≈ 1 y^{-1}

Green area: 90% C.L. region calculated taking into account the signal in (2-6) keV

Clear annual modulation in (1-6) keV single-hit scintillation events

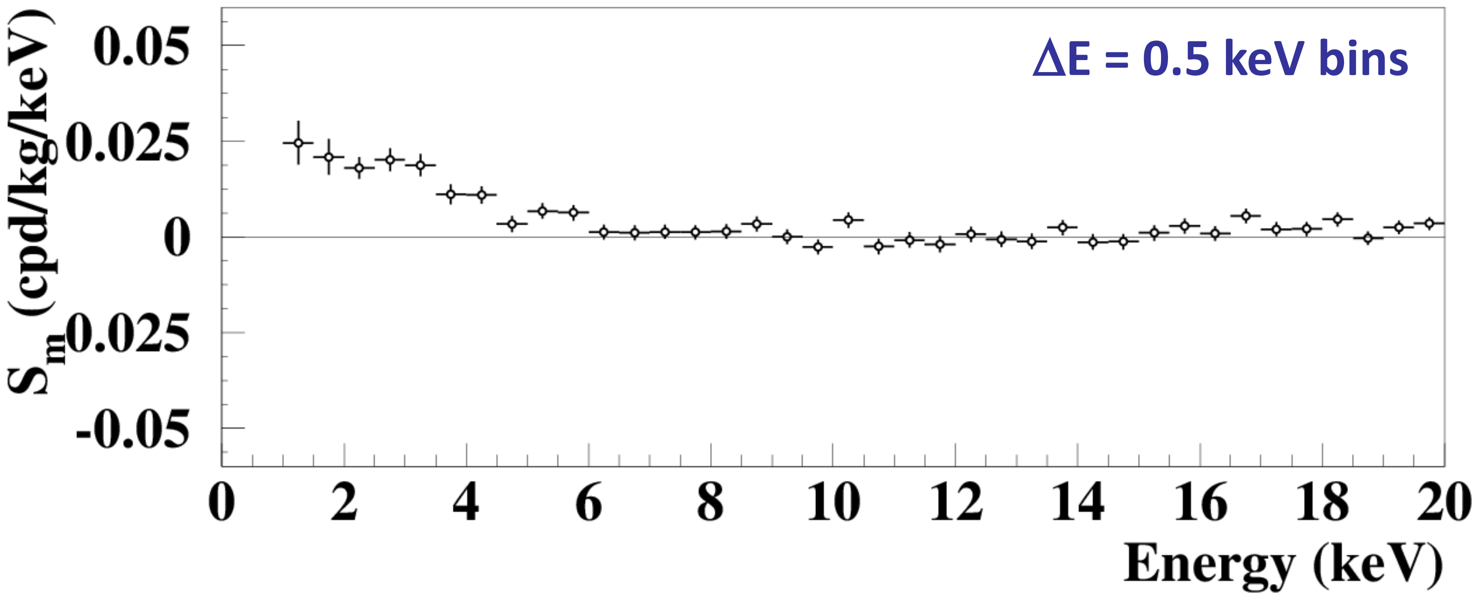
Energy distribution of the modulation amplitudes

Max-likelihood analysis

$$R(t) = S_0 + S_m \cos[\omega(t - t_0)]$$

here $T=2\pi/\omega=1$ yr and $t_0= 152.5$ day

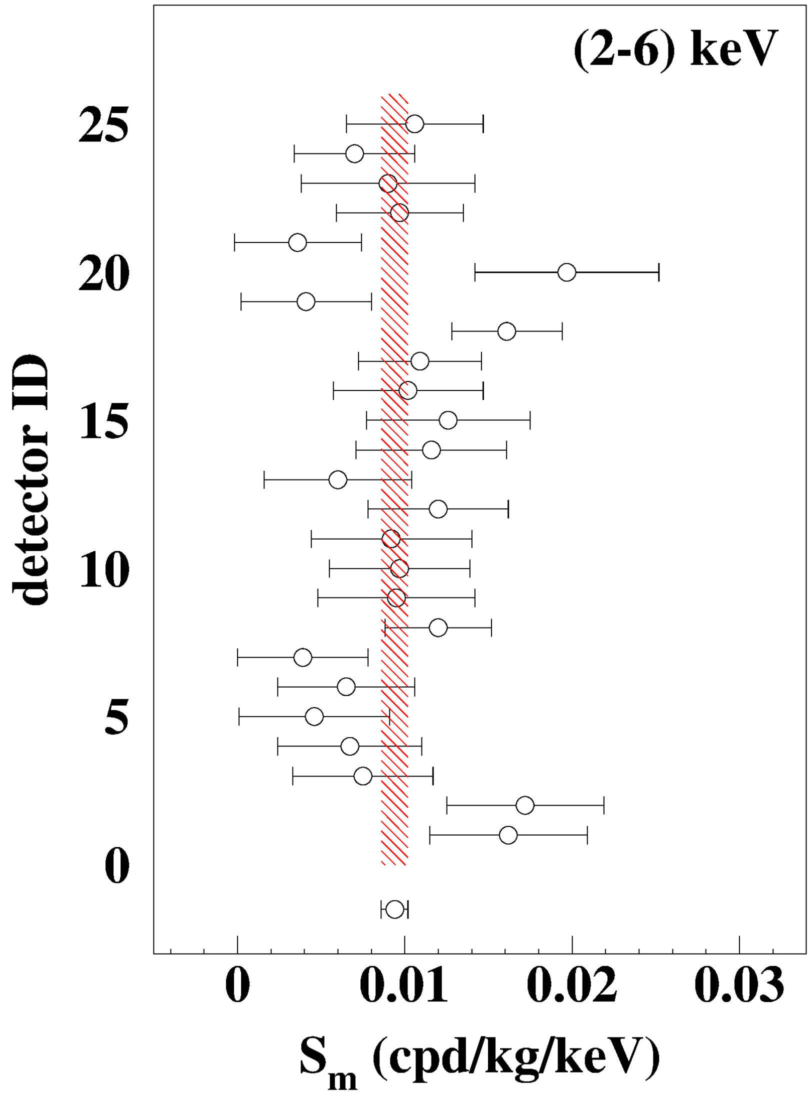
DAMA/NaI + DAMA/LIBRA-phase1
+ DAMA/LIBRA-phase2 (2.46 ton×yr)



A clear modulation is present in the (1-6) keV energy interval, while S_m values compatible with zero are present just above

- The S_m values in the (6–14) keV energy interval have random fluctuations around zero with χ^2 equal to 19.0 for 16 degrees of freedom (upper tail probability 27%).
- In (6–20) keV χ^2/dof = 42.6/28 (upper tail probability 4%). The obtained χ^2 value is rather large due mainly to two data points, whose centroids are at 16.75 and 18.25 keV, far away from the (1–6) keV energy interval. The P-values obtained by excluding only the first and either the points are 11% and 25%.

S_m for each detector



DAMA/LIBRA-phase1 +
DAMA/LIBRA-phase2
total exposure: **2.17 ton×yr**

S_m integrated in the range (2 - 6) keV for each of the 25 detectors (1σ error)

Shaded band = weighted averaged $S_m \pm 1\sigma$

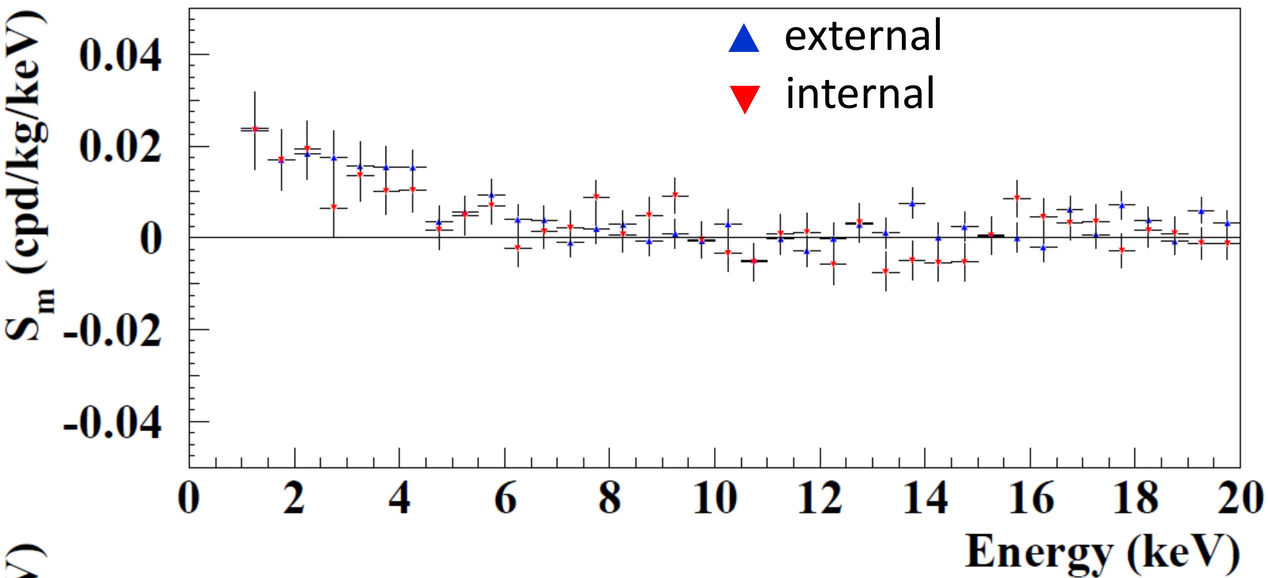
χ^2/dof = 23.9/24 d.o.f.

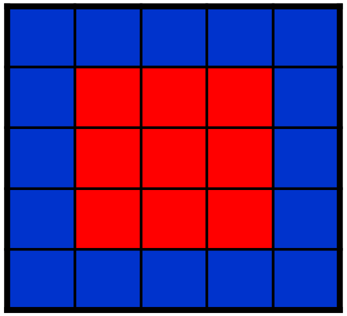
The signal is well distributed over all the 25 detectors

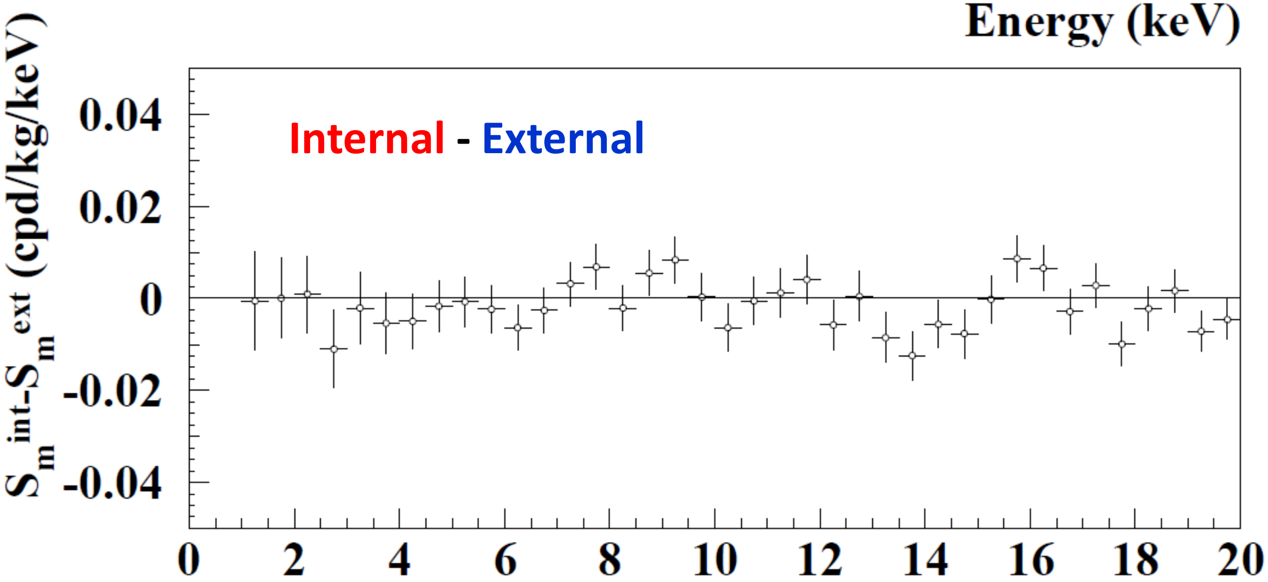
External vs internal detectors

DAMA/LIBRA-phase2

ΔE=0.5 keV







1-4 keV χ^2**/dof =2.5/6**

1-10 keV χ^2**/dof =12.1/8**

1-20 keV χ^2**/dof =40.8/38**

- Contributions to the total neutron flux at LNGS;
- Counting rate in DAMA/LIBRA for *single-hit* events, in the (2 – 6) keV energy region induced by:
  - neutrons,
  - muons,
  - solar neutrinos.

$$\Phi_k = \Phi_{0,k}\left(1 + \eta_k cos\omega\left(t - t_k\right)\right)$$
$$R_k = R_{0,k}\left(1 + \eta_k cos\omega\left(t - t_k\right)\right)$$

EPJC 74 (2014) 3196 (also EPJC 56 (2008) 333, EPJC 72 (2012) 2064, IJMPA 28 (2013) 1330022)

Modulation amplitudes

Source		$\Phi_{0,k}^{(n)}$ (neutrons cm^{-2} s^{-1})	η_k	t_k	$R_{0,k}$ (cpd/kg/keV)		$A_k = R_{0,k}\eta_k$ (cpd/kg/keV)	A_k/S_m^{exp}
SLOW neutrons	thermal n ($10^{-2}-10^{-1}$ eV)	1.08×10^{-6} [15]	$\simeq 0$ however $\ll 0.1$ [2, 7, 8]	–	$< 8 \times 10^{-6}$	[2, 7, 8]	$\ll 8 \times 10^{-7}$	$\ll 7 \times 10^{-5}$
	epithermal n (eV-keV)	2×10^{-6} [15]	$\simeq 0$ however $\ll 0.1$ [2, 7, 8]	–	$< 3 \times 10^{-3}$	[2, 7, 8]	$\ll 3 \times 10^{-4}$	$\ll 0.03$
FAST neutrons	fission, $(\alpha, n) \to$ n (1-10 MeV)	$\simeq 0.9 \times 10^{-7}$ [17]	$\simeq 0$ however $\ll 0.1$ [2, 7, 8]	–	$< 6 \times 10^{-4}$	[2, 7, 8]	$\ll 6 \times 10^{-5}$	$\ll 5 \times 10^{-3}$
	$\mu \to$ n from rock ($>$ 10 MeV)	$\simeq 3 \times 10^{-9}$ (see text and ref. [12])	0.0129 [23]	end of June [23, 7, 8]	$\ll 7 \times 10^{-4}$	(see text and [2, 7, 8])	$\ll 9 \times 10^{-6}$	$\ll 8 \times 10^{-4}$
	$\mu \to$ n from Pb shield ($>$ 10 MeV)	$\simeq 6 \times 10^{-9}$ (see footnote 3)	0.0129 [23]	end of June [23, 7, 8]	$\ll 1.4 \times 10^{-3}$	(see text and footnote 3)	$\ll 2 \times 10^{-5}$	$\ll 1.6 \times 10^{-3}$
	$\nu \to$ n (few MeV)	$\simeq 3 \times 10^{-10}$ (see text)	0.03342 *	Jan. 4th *	$\ll 7 \times 10^{-5}$	(see text)	$\ll 2 \times 10^{-6}$	$\ll 2 \times 10^{-4}$
	direct μ	$\Phi_0^{(\mu)} \simeq 20\ \mu$ m^{-2}d^{-1} [20]	0.0129 [23]	end of June [23, 7, 8]	$\simeq 10^{-7}$	[2, 7, 8]	$\simeq 10^{-9}$	$\simeq 10^{-7}$
	direct ν	$\Phi_0^{(\nu)} \simeq 6 \times 10^{10}\ \nu$ cm^{-2}s^{-1} [26]	0.03342 *	Jan. 4th *	$\simeq 10^{-5}$	[31]	3×10^{-7}	3×10^{-5}

* The annual modulation of solar neutrino is due to the different Sun-Earth distance along the year; so the relative modulation amplitude is twice the eccentricity of the Earth orbit and the phase is given by the perihelion.

All are negligible w.r.t. the annual modulation amplitude observed by DAMA/LIBRA and they cannot contribute to the observed modulation amplitude.

+ In no case neutrons (of whatever origin) can mimic the DM annual modulation signature since some of the **peculiar requirements of the signature** would fail, such as the neutrons would induce e.g. variations in all the energy spectrum, variation in the multiple hit events,… which were not observed.

Summary of the results obtained in the additional investigations of possible systematics or side reactions – DAMA/LIBRA

NIMA592(2008)297, EPJC56(2008)333, J. Phys. Conf. ser. 203(2010)012040, arXiv:0912.0660, S.I.F.Atti Conf.103(211), Can. J. Phys. 89 (2011) 11, Phys.Proc.37(2012)1095, EPJC72(2012)2064, arxiv:1210.6199 & 1211.6346, IJMPA28(2013)1330022, EPJC74(2014)3196, IJMPA31(2017)issue31, Universe4(2018)03009, Beld19,2(2018)27

Source	Main comment	Cautious upper limit (90%C.L.)
RADON	Sealed Cu box in HP Nitrogen atmosphere, 3-level of sealing, etc.	**$<2.5\times10^{-6}$ cpd/kg/keV**
TEMPERATURE	Installation is air conditioned+ detectors in Cu housings directly in contact with multi-ton shield→ huge heat capacity + T continuously recorded	**$<10^{-4}$ cpd/kg/keV**
NOISE	Effective full noise rejection near threshold	**$<10^{-4}$ cpd/kg/keV**
ENERGY SCALE	Routine + intrinsic calibrations	**$<1\text{-}2\times10^{-4}$ cpd/kg/keV**
EFFICIENCIES	Regularly measured by dedicated calibrations	**$<10^{-4}$ cpd/kg/keV**
BACKGROUND	No modulation above 6 keV; no modulation in the (2-6) keV *multiple-hits* events; this limit includes all possible sources of background	**$<10^{-4}$ cpd/kg/keV**
SIDE REACTIONS	Muon flux variation measured at LNGS	**$<3\times10^{-5}$ cpd/kg/keV**

+ they cannot satisfy all the requirements of annual modulation signature

Thus, they cannot mimic the observed annual modulation effect

Final model independent result
DAMA/NaI+DAMA/LIBRA-phase1+phase2

Presence of modulation over 20 annual cycles at 12.9 σ C.L. with the proper distinctive features of the DM signature; all the features satisfied by the data over 20 independent experiments of 1 year each one

The total exposure by former DAMA/NaI, DAMA/LIBRA-phase1 and phase2 is 2.46 ton × yr

In fact, as required by the DM annual modulation signature:

1) **The *single-hit* events show a clear cosine-like modulation, as expected for the DM signal**

2) **Measured period is equal to (0.999±0.001)* yr, well compatible with the 1 yr period, as expected for the DM signal**

3) **Measured phase (145±5)* days is well compatible with the roughly about 152.5 days as expected for the DM signal**

4) **The modulation is present only in the low energy (2—6) keV energy interval and not in other higher energy regions, consistently with expectation for the DM signal**

5) **The modulation is present only in the *single-hit* events, while it is absent in the *multiple-hit* ones as expected for the DM signal**

6) **The measured modulation amplitude in NaI(Tl) of the *single-hit* events is: (0.0103 ± 0.0008)* cpd/kg/keV (12.9 σ C.L.).**

* Here 2-6 keV energy interval

No systematic or side process able to simultaneously satisfy all the many peculiarities of the signature and to account for the whole measured modulation amplitude is available

… and well compatible with several candidates
(in many possible astrophysical, nuclear and particle physics scenarios)

Examples of uncertainties in models and scenarios

see for some details e.g.:
Riv.N.Cim.26 n.1 (2003) 1, IJMPD13(2004)2127, EPJC47 (2006)263, IJMPA21 (2006)1445

Nature of the candidate and couplings

- WIMP class particles (neutrino, sneutrino, etc.): SI, SD, mixed SI&SD, preferred inelastic + e.m. contribution in the detection
- Light bosonic particles
- Kaluza-Klein particles
- Mirror dark matter
- Heavy Exotic candidate
- …etc. etc.

Halo models & Astrophysical scenario

- Isothermal sphere ⇒ very simple but unphysical halo model
- Many consistent halo models with different density and velocity distribution profiles can be considered with their own specific parameters (see e.g. PRD61(2000)023512)
- Caustic halo model
- Presence of non-thermalized DM particle components
- Streams due e.g. to satellite galaxies of the Milky Way (such as the Sagittarius Dwarf)
- Multi-component DM halo
- Clumpiness at small or large scale
- Solar Wakes
- …etc. …

Instrumental quantities

- Energy resolution
- Efficiencies
- Quenching factors
- Channeling effects
- Their dependence on energy
- …

Quenching Factor

- differences are present in different experimental determinations of q for the same nuclei in the same kind of detector depending on its specific features (e.g. q depends on dopant and on the impurities; in liquid noble gas e.g.on trace impurities, on presence of degassing/releasing materials, on thermodynamical conditions, on possibly applied electric field, etc); assumed 1 in bolometers
- channeling effects possible increase at low energy in scintillators (dL/dx)
- possible larger values of q (AstropPhys33 (2010) 40)

→ energy dependence

Scaling laws
of cross sections for the case of recoiling nuclei

- Different scaling laws for different DM particle:

$\sigma_A \propto \mu^2 A^2 (1+\varepsilon_A)$

$\varepsilon_A = 0$ *generally assumed*

$\varepsilon_A \approx \pm 1$ *in some nuclei? even* for neutralino candidate in MSSM (see Prezeau, Kamionkowski, Vogel et al., PRL91(2003)231301)

Form Factors
for the case of recoiling nuclei

- Many different profiles available in literature for each isotope
- Parameters to fix for the considered profiles
- Dependence on particle-nucleus interaction
- In SD form factors: no decoupling between nuclear and Dark Matter particles degrees of freedom + dependence on nuclear potential

Spin Factors
for the case of recoiling nuclei

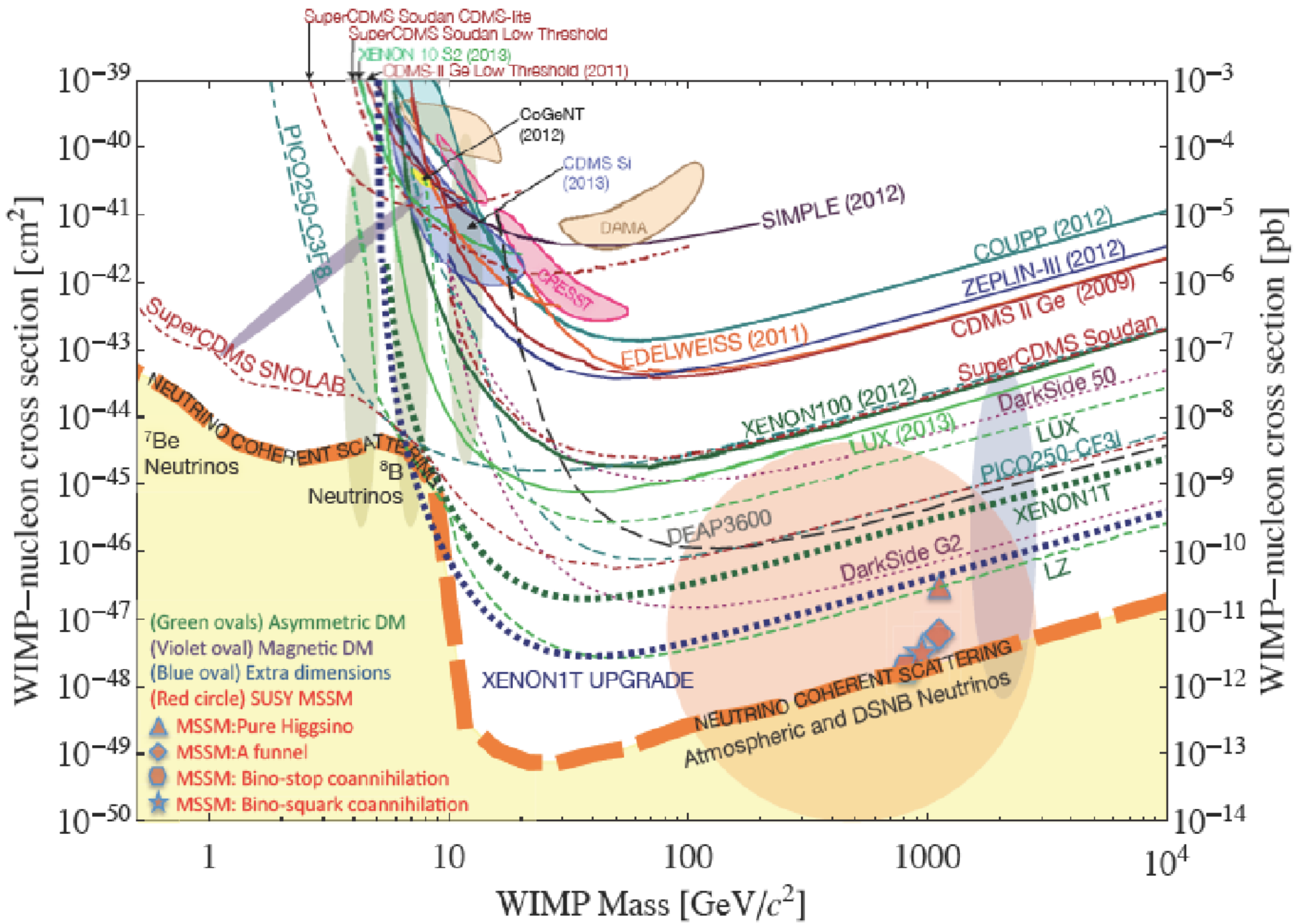
- Calculations in different models give very different values also for the same isotope
- Depend on the nuclear potential models
- Large differences in the measured counting rate can be expected using:

  either SD not-sensitive isotopes

  or SD sensitive isotopes depending on the unpaired nucleon (compare e.g. odd spin isotopes of Xe, Te, Ge, Si, W with the ^{23}Na and ^{127}I cases).

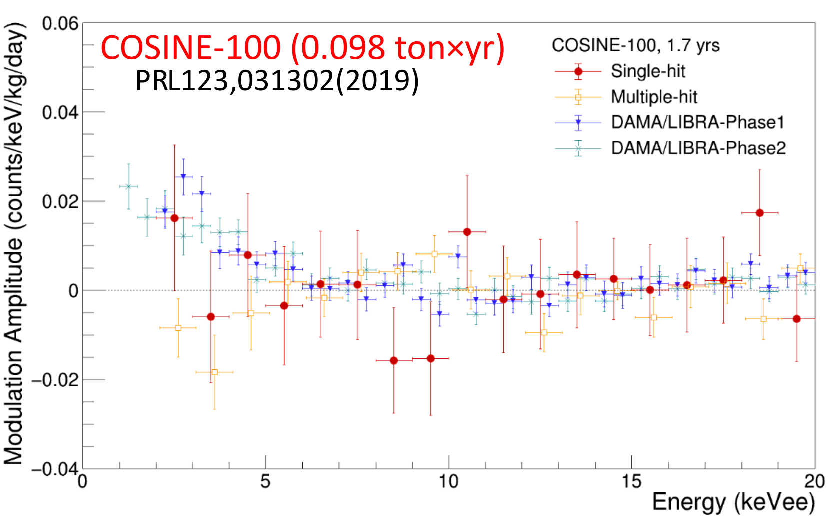
… and more …

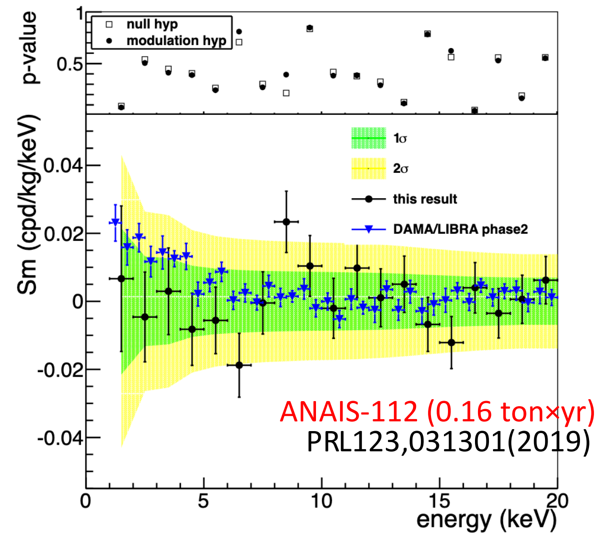
Is it an “universal” and “correct” way to approach the problem of DM and comparisons?



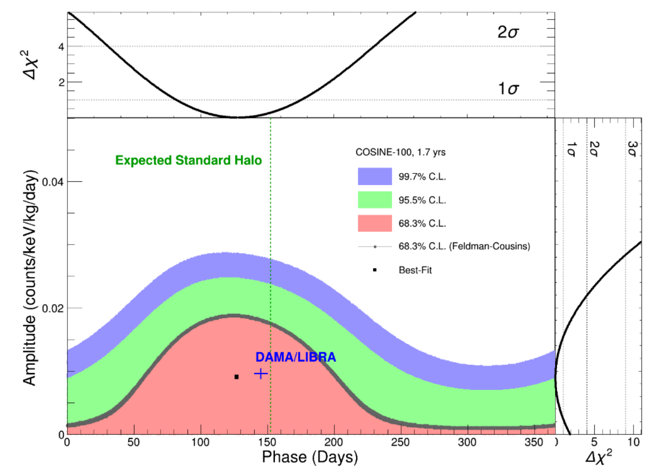
No, it isn’t. This is just a largely arbitrary/partial/incorrect exercise

Other annual modulation results with other NaI(Tl)





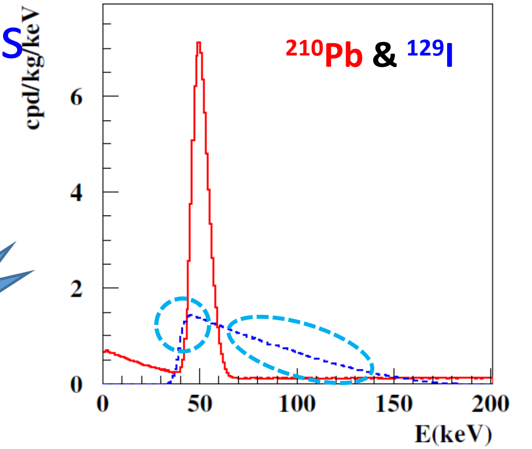
Energy interval	Experiment	Exposure ton x yr	Rate (cpd/kg/keV)	Amplitude (cpd/kg/keV)
(2,6) keV	DAMA/LIBRA (ph1 + ph2)	2.17	0.8	0.0095 ± 0.0008
	COSINE-100	0.098	3.0	0.0083 ± 0.0068
	ANAIS-112	0.16	3.2	- 0.0044 ± 0.0058
(1,6) keV	DAMA/LIBRA-phase2	1.13	0.7	0.0105 ± 0.0011
	ANAIS-112	0.16	3.6	- 0.0015 ± 0.0063



DAMA-LIBRA is still much better than any other NaI(Tl) experiment for exposure time, for exposed mass, for background, for energy threshold and control of all the experimental parameters.

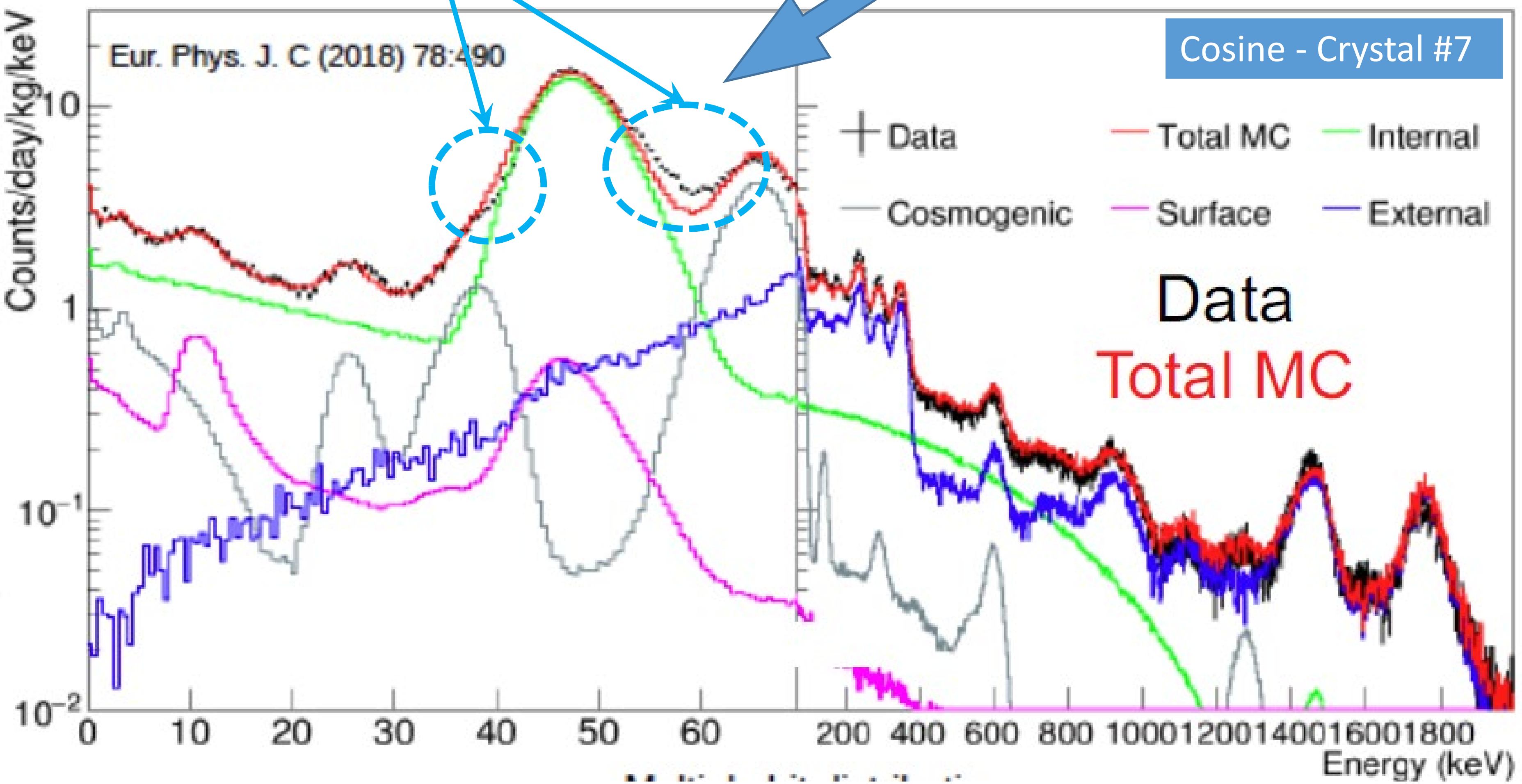
COSINE & ANAIS data have not sufficient sensitivity to DAMA signal

^{129}I completely forgotten in Cosine-100 data analysis



Very important discrepancies (note the log scale) in the reconstruction of the structure at ≈ 45 keV, due to:

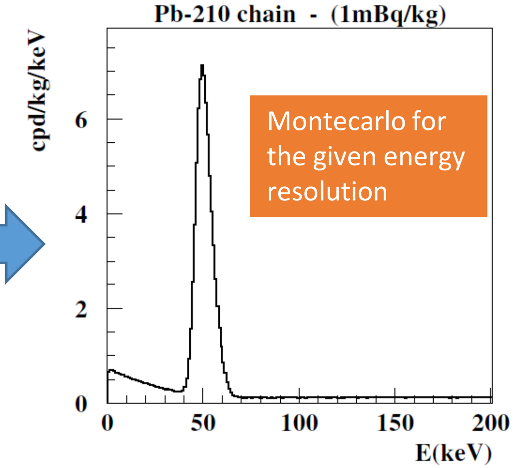
1. **Missing contribute of ^{129}I**
2. **Overestimate contribute of ^{210}Pb**

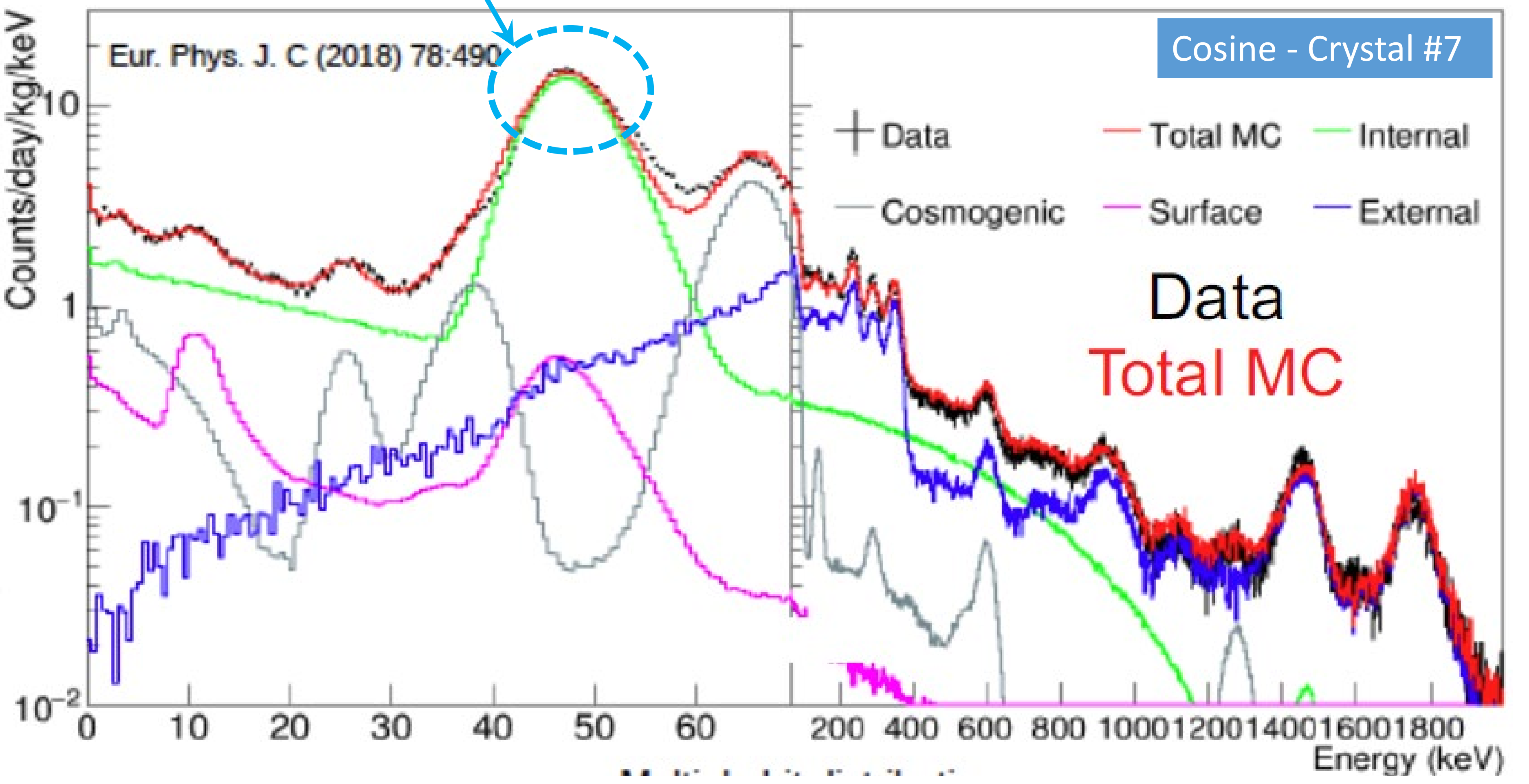


… and ^{210}Pb significantly overestimated

In green spectrum, the ^{210}Pb peak height is ≈ 14cpd/kg/keV, that is ≈ 2mBq/kg

But the measured α rate in crystal 7 is (1.54±0.4) mBq/kg and this should be an upper limit for ^{210}Pb activity!





In conclusion:

the Cosine-100 low energy analysis is wrong and the exclusion plot meaningless

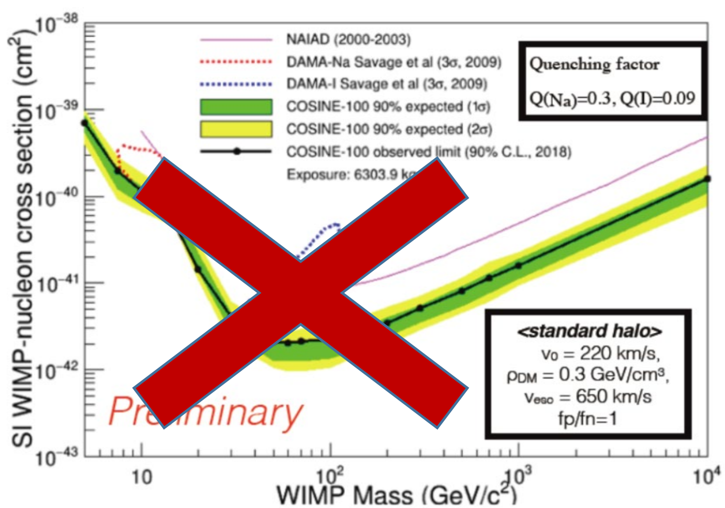
Cosine - Crystal #7

Components	Background 2-6 keV (dru)
Internal ^{210}Pb	1.50 +/- 0.07
Internal ^{40}K	0.05 +/- 0.01
Surface ^{210}Pb	0.38 +/- 0.21
^{3}H (Cosmogenic)	0.58 +/- 0.54
^{109}Cd (Cosmogenic)	0.09 +/- 0.09
Other cosmogenic	0.05 +/-0.03
External	0.03 +/- 0.02
Total expected	**2.70 +/- 0.59**
Data	**2.64 +/- 0.05**

Internal ^{210}Pb seems to give the main (≈60%) contribution in 2-6 keV region, but , as shown, the assumed value is wrong: < 1.2 dru

To be revised

Wrong: expected << observed
Large space for DM signal



An example: how not to do to get a result (exclusion limits)
The case of COSINE-100

- The methodology of the background subtraction, used for example by Cosine-100, is strongly discouraged and deprecated because of the impossibility to have a precise knowledge of the background contribution in particular at low energy, leading to large systematic uncertainties.

Very important discrepancies in the reconstruction of the structure at ≈ 45 keV, due to:

1. Missing contribute of ^{129}I (emended in a later paper, but not in the exclusion limits))
2. Overestimate contribute of ^{210}Pb

Components	Background 2-6 keV (dru)
Internal ^{210}Pb	1.50 +/- 0.07
Internal ^{40}K	0.05 +/- 0.01
Surface ^{210}Pb	0.38 +/- 0.21
^{3}H (Cosmogenic)	0.58 +/- 0.54
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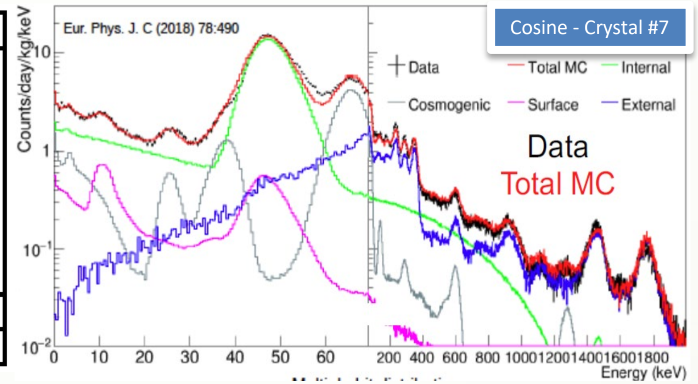


- ✓ Even considering the background model as correct, the analysis has fault.
- ✓ They get null residuals in each crystal (even always negative) starting from a wrong bckg hypothesis!

Data−model = −0.105±0.276 cpd/kg/keV

→ S_0<0.36 cpd/kg/keV 90%CL in the (2-6) keV energy region

Still large space for DM

Since time, by simple and direct determination in DAMA: S_0<0.18 cpd/kg/keV in (2-4) keV (DAMA/LIBRA-phase2).

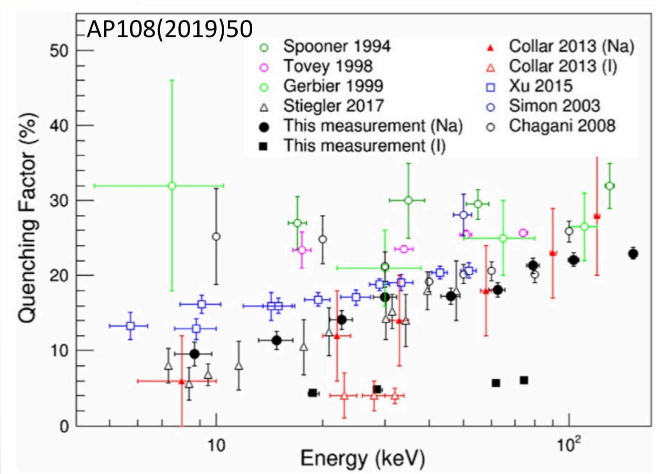
Cosine-100 low energy analysis is wrong and the exclusion limits are meaningless (published on Nature!!)

In conclusion: the methodology of the background subtraction is a **dangerous** way to claim sensitivities by the fact not supported by large counting rate

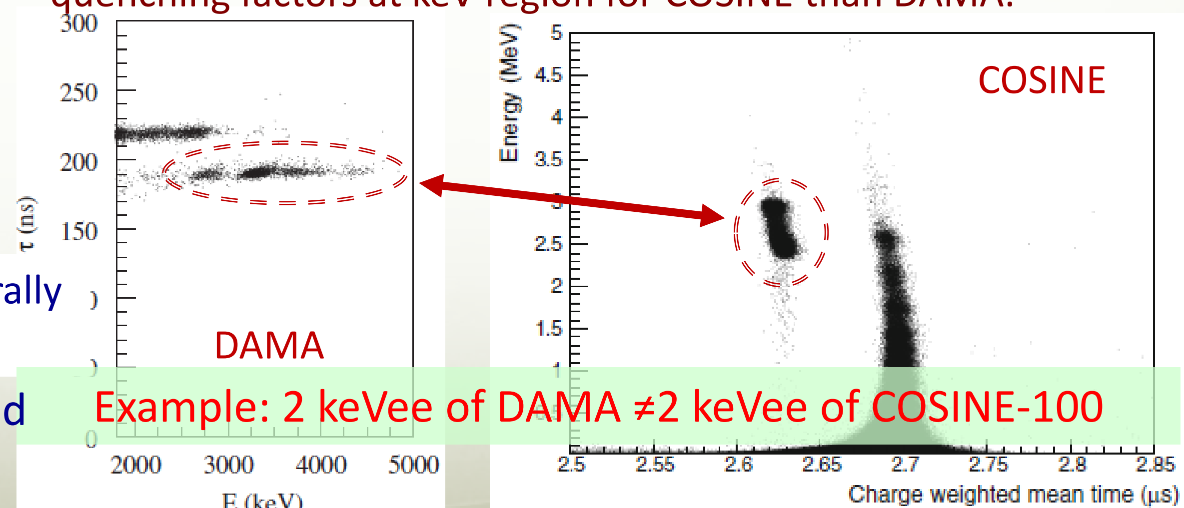
The case of the NaI(Tl) quenching factors (QF)

- ✓ The QFs are a property of the specific detector and not general property, particularly in the very low energy range.
- ✓ For example in NaI(Tl), QFs depend on the adopted growing procedures, on Tl concentration and uniformity in the detector, on the specific materials added in the growth, on the mono-crystalline or poly-crystalline nature of the detector, etc.
- ✓ Their measurements are difficult and always affected by significant experimental uncertainties.
- ✓ All these aspects are always relevant sources of uncertainties when comparing whatever results in terms of DM candidates inducing nuclear recoils.

+ QF depending on energy + channeling effects
+ Migdal effect



- A wide spread existing in literature for NaI(Tl)
- This is also confirmed by the different α/β light ratio measured with DAMA and COSINE crystals. This implies much lower quenching factors at keV region for COSINE than DAMA.



CURIOSITY: Recent productions (generally by Bridgman growth) yields low QF…

Example: 2 keVee of DAMA ≠2 keVee of COSINE-100

The model dependent analyses and comparisons must be performed using the QF **measured** for each detector.

Alphas from ^{238}U and ^{232}Th chains span from 2.6 to 4.5 MeVee in DAMA, while from 2.3 to 3.0 MeVee in COSINE

Examples of model-dependent analyses

DM particles elastically interacting with target nuclei – SI interaction

DAMA/NaI, DAMA/LIBRA-ph1 and ph2

arXiv:1907.06405

- A large (but not exhaustive) class of halo models is considered;
- Local velocity v_0 in the range [170,270] km/s;
- Halo density ρ depending on the halo model;
- v_{esc} = 550 km/s (no sizable differences if v_{esc} in the range [550, 650]km/s);
- For DM candidates inducing nuclear recoils: three different sets of values for the nuclear form factor and quenching factor parameters.

σ_{SI} SI point-like DM-nucleon cross section

ξ fractional amount of local density in terms of the considered DM candidate

The point-like SI cross section of DM particles scattering off (A,Z) nucleus:

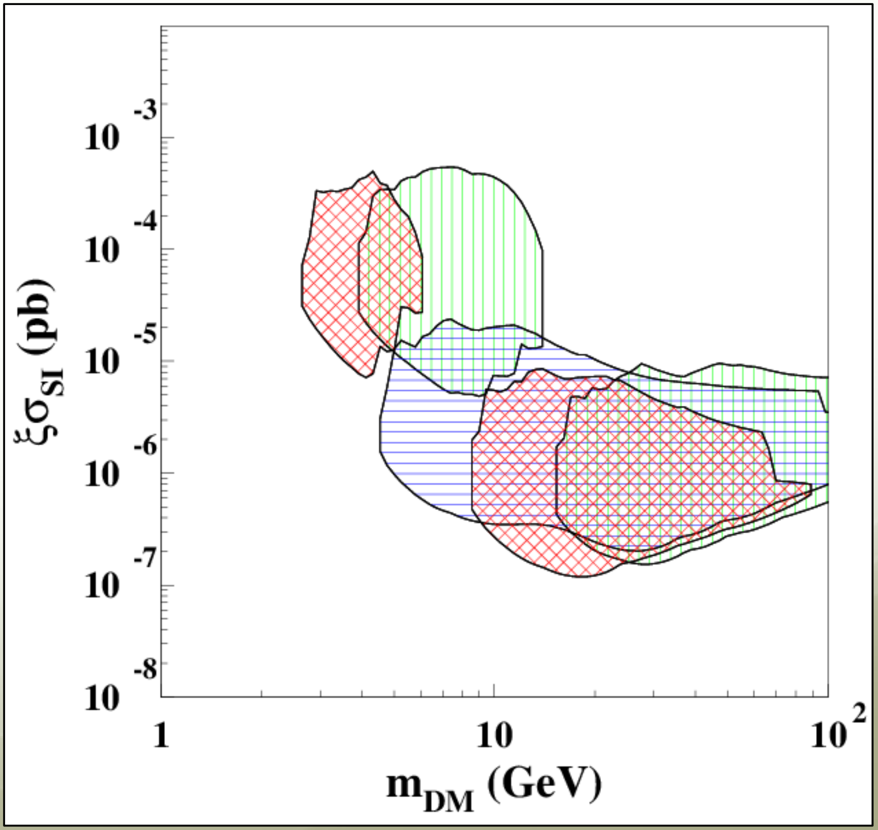
$$\sigma_{SI}(A,Z) \propto m^2_{red}(A,DM)\left[f_p Z + f_n (A-Z)\right]^2$$

where f_p, f_n are the effective DM particle couplings to protons and neutrons.

If $f_p=f_n$: $$\sigma_{SI}(A,Z) = \frac{m^2_{red}(A,DM)}{m^2_{red}(1,DM)} A^2 \sigma_{SI}$$

$\xi\sigma_{SI}$ vs m_{DM}

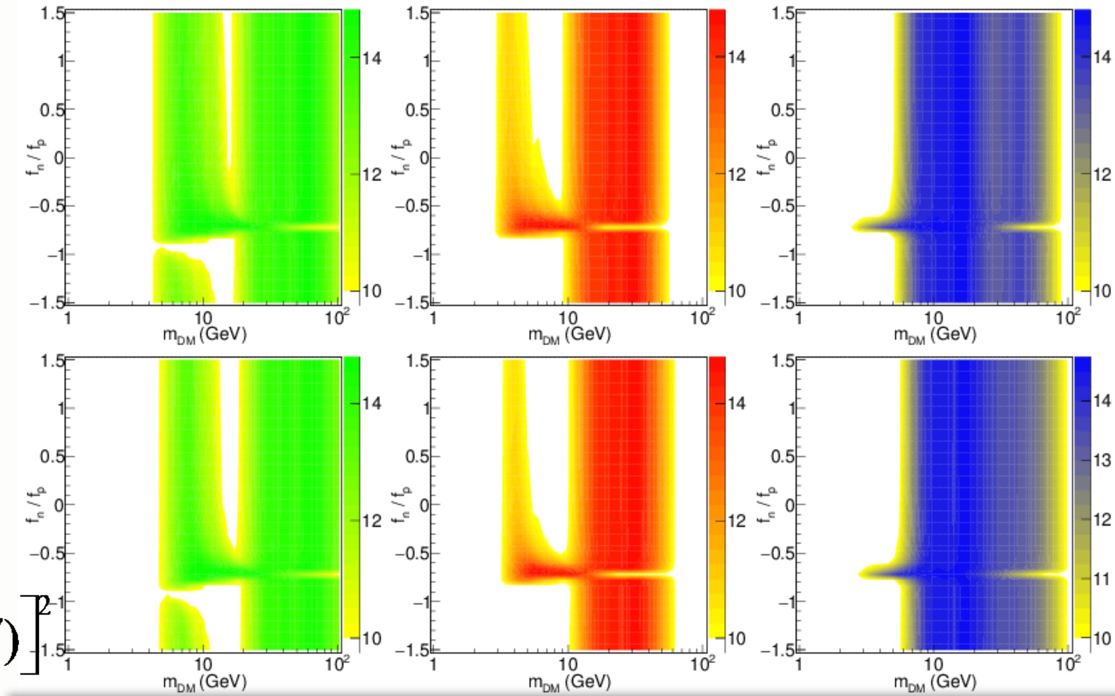
1. Constants q.f.
2. Varying q.f.(E_R)
3. With channeling effect



Allowed DAMA regions:
Domains where the likelihood-function values differ more than 10σ from absence of signal

Model-dependent analyses

DM particles elastically interacting with target nuclei SI-IV interaction



DAMA/NaI, DAMA/LIBRA-ph1 and ph2

Case of isospin violating SI coupling:

$$f_p \neq f_n$$

$$\sigma_{SI}(A,Z) \propto m_{red}^2(A,DM)\left[f_p Z + f_n(A-Z)\right]^2$$

f_n/f_p vs m_{DM}
marginalizing on $\xi\sigma_{SI}$

1. **Constants q.f.**
2. **Varying q.f.(E_R)**
3. **With channeling effect**

Allowed DAMA regions for A0 (isothermal sphere), B1, C1, D3 halo models (top to bottom)

- Two bands at low mass and at higher mass;
- Good fit for low mass DM candidates at $f_n/f_p \approx$ -53/74 = = -0.72 (signal mostly due to ^{23}Na recoils).
- Contrary to what was stated in Ref. [PLB789,262(2019), JCAP07,016(2018), JCAP05,074(2018)] where the low mass DM candidates were disfavored for f_n/f_p = 1 by DAMA data, the inclusion of the uncertainties related to halo models, quenching factors, channeling effect, nuclear form factors, etc., can also support low mass DM candidates either including or not the channeling effect.
- The case of isospin-conserving f_n/f_p=1 is well supported at different extent both at lower and larger mass.

Model-dependent analyses: other examples

arXiv:1907.06405

DM particles elastically interacting with target nuclei – purely SD interaction

Only possible for target nuclei with spin≠0 $\tan\vartheta = \frac{a_n}{a_p}, \quad \vartheta \text{ in } [0,\pi]$

a_p and a_n are the effective DM-nucleon coupling strengths for SD int.

$\theta = 0 \Rightarrow a_n=0,\ a_p \neq 0$ or $|a_p| >> |a_n|$;

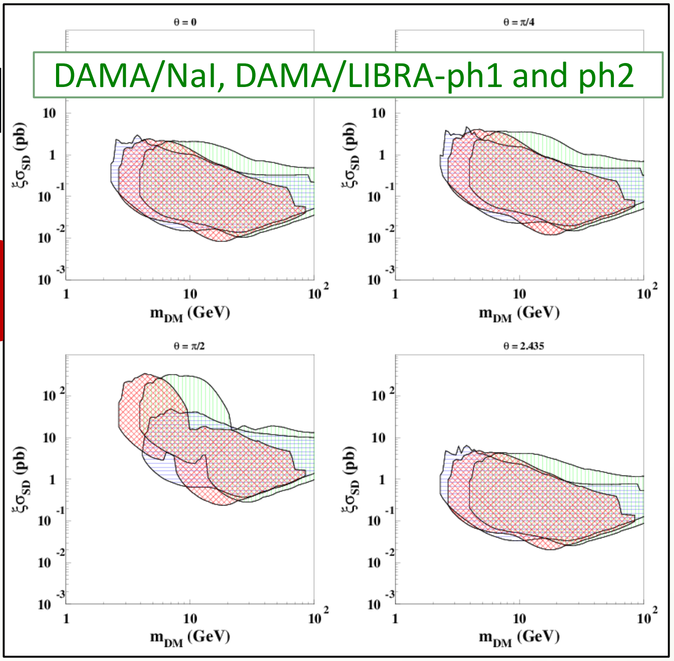
$\theta = \pi/4 \Rightarrow a_n = a_p$;

$\theta = \pi/2 \Rightarrow a_p=0,\ a_n \neq 0$ or $|a_n| >> |a_p|$;

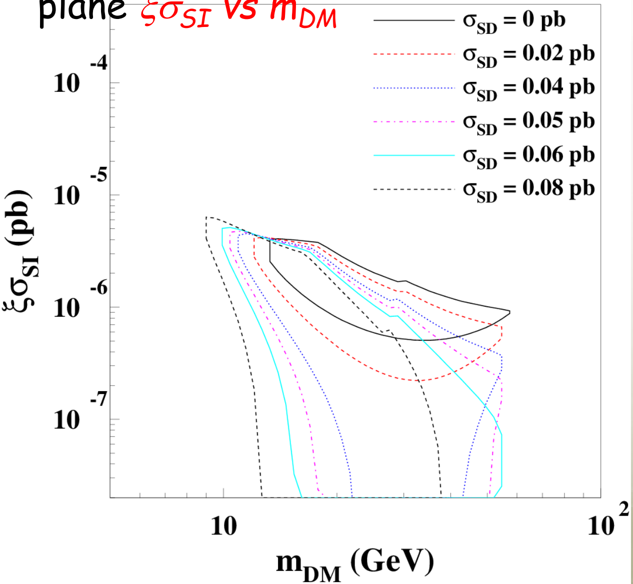
$\theta = 2.435\text{rad} \Rightarrow a_n/a_p = -0.85$, pure Z_0 coupling

$\xi\sigma_{SD}$ vs m_{DM}

1. Constants q.f.
2. Varying q.f.(E_R)
3. With channeling effect



Effect induced by the inclusion of a SD component on allowed regions in the plane $\xi\sigma_{SI}$ vs m_{DM}



- Even a relatively small SD (SI) contribution can drastically change the allowed region in the (m_{DM}, $\xi\sigma_{SI(SD)}$) plane;
- The model-dependent comparison plots between exclusion limits at a given C.L. and regions of allowed parameter space do not hold e.g. for mixed scenarios when comparing experiments with and without sensitivity to the SD component of the interaction.
- The same happens when comparing regions allowed by experiments whose target-nuclei have unpaired proton with exclusion plots quoted by experiments using target-nuclei with unpaired neutron when the SD component of the interaction would correspond either to $\theta \approx 0$ or $\theta \approx \pi$

DAMA/LIBRA towards the lowering of the software energy threshold

Model-independent evidence by DAMA/NaI and DAMA/LIBRA

well compatible with several candidates in many astrophysical, nuclear and particle physics scenarios

Just few examples of interpretation of the annual modulation in terms of candidate particles in some scenarios

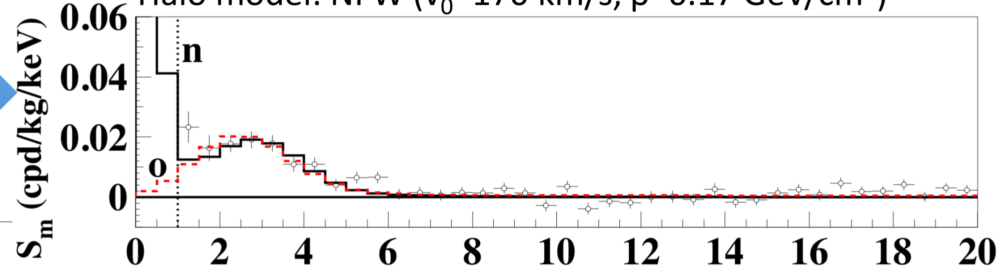
LDM candidates

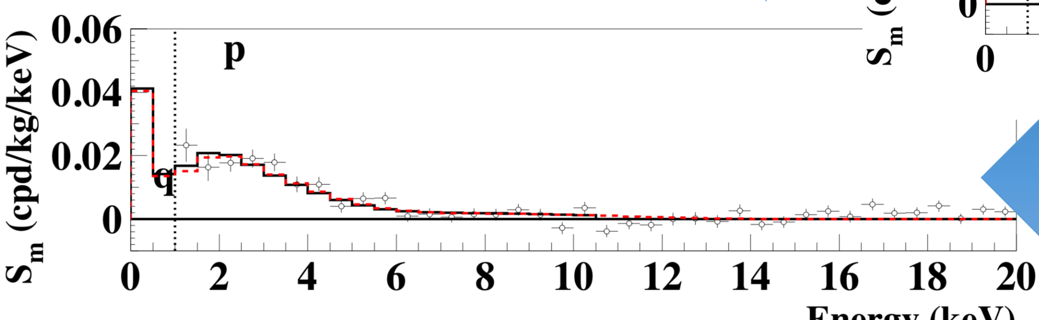
Halo model: NFW (v_0=170 km/s, ρ=0.17 GeV/cm[3])

LDM with coherent scattering on nuclei

n - m_H=30 MeV, δ=16 MeV σ=1.4x10^{-6} pb

o - m_H=100 MeV, δ=45 MeV σ=2.4x10^{-6} pb





LDM with incoherent scattering on nuclei

p - m_H=30 MeV, δ=8 MeV σ=0.008 pb

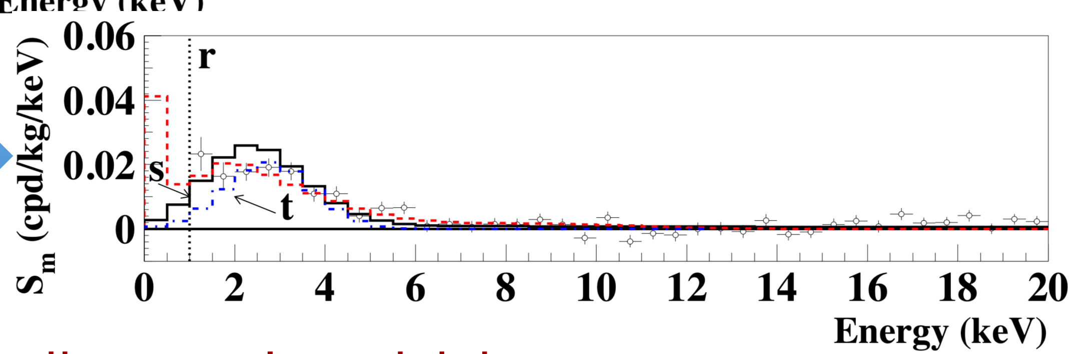
q - m_H=100 MeV, δ=2 MeV σ=0.027 pb

LDM with m_L=0 GeV (δ=m_H)

r - coherent on nucl. m_H=28 MeV, σ=1.3x10^{-6} pb

s - incoherent on nucl. m_H=20 MeV, σ=0.006 pb

t – on electrons m_H=56 keV, σ=2.3x10^{-7} pb

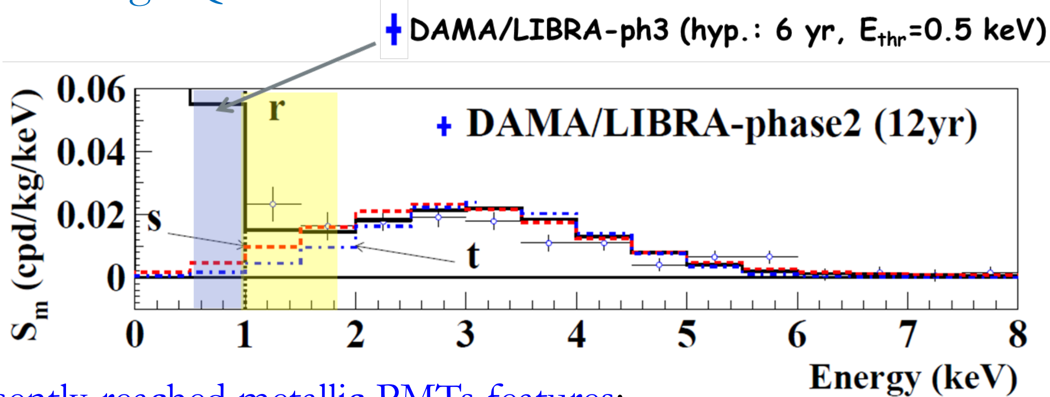


Compatibility with several candidates;

other ones are open

Lowering the software energy threshold below 1 keV with high overall efficiency

new miniaturized low background **pre-amps** directly installed on the low-background supports of the **voltage dividers** of the new lower background high Q.E. **PMTs**

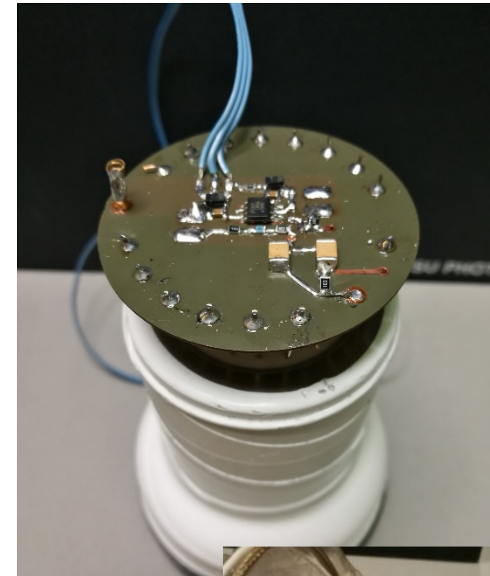


The presently-reached metallic PMTs features:

- Q.E. around 35-40% @ 420 nm (NaI(Tl) light)
- Radio-purity at level of 5 mBq/PMT (^{40}K), 3-4 mBq/PMT (^{232}Th), 3-4 mBq/PMT (^{238}U), 1 mBq/PMT (^{226}Ra), 2 mBq/PMT (^{60}Co).
- Dark counts < 100 Hz

The features of the voltage divider+preamp system:
S/N improvement ≈3.0-9.0, discrimination of the single ph.el. from electronic noise: 3 – 8, the Peak/Valley ratio: 4.7 - 11.6; residual radioactivity much lower than that of the single PMT.

(If the tests will be satisfactory we plan to replace all PMTs, otherwise the electronics (TD + voltage divider + preamp.) upgrade is planned)



- several prototypes from a dedicated R&D with HAMAMATSU at hand
- 4 DAMA/LIBRA detectors equipped with the new PMTs as required by CSN2 referees

Anisotropic scintillators

Measurements of $ZnWO_4$ anisotropic response to nuclear recoils for the ADAMO project

[among the DAMA activities from June 2019 to June 2020]

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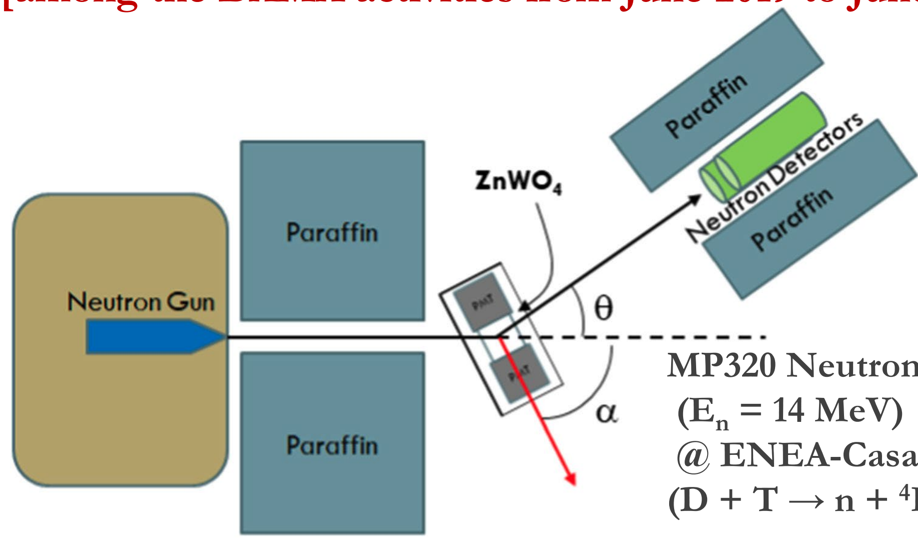


MP320 Neutron Generator
(E_n = 14 MeV)
@ ENEA-Casaccia lab
($D + T \rightarrow n + {}^4He$)

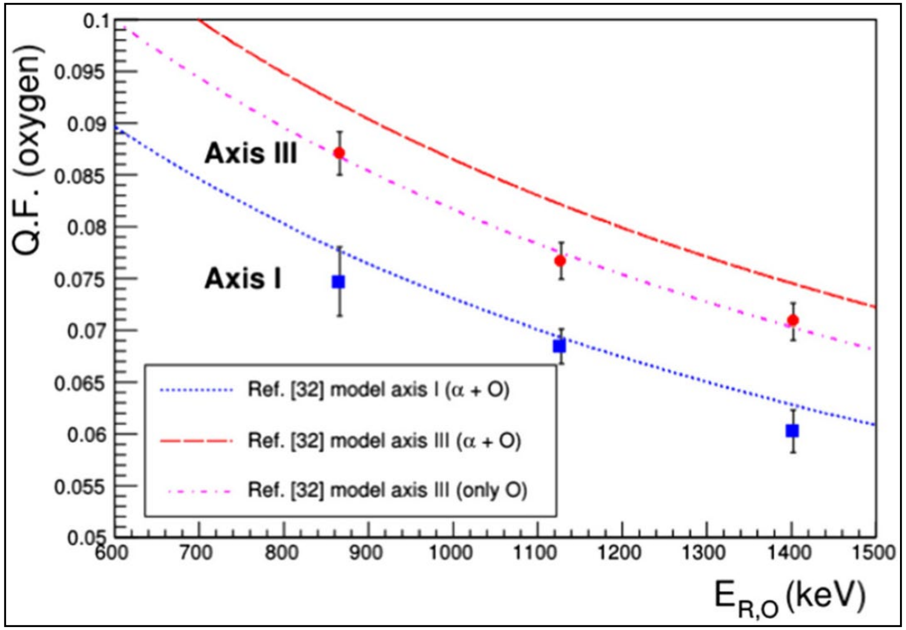
Anisotropic scintillators can offer a unique possibility to exploit the directionality approach in order to investigate the presence of those Dark Matter candidates inducing just nuclear recoils

A neutron generator at ENEA-CASACCIA Lab. and neutron detectors to tag the scattered neutrons have been used to measure the anisotropic response to nuclear recoils

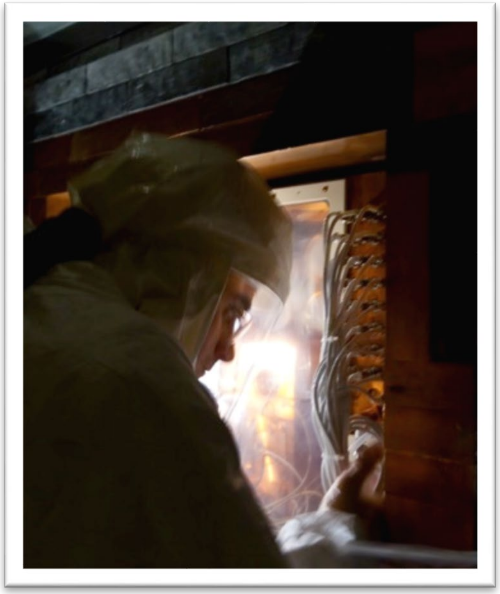
Measure of quenching factors for nuclear recoils for different crystallographic axes and nuclear recoils energies

First measurement of anisotropy for recoils in energy region down to some hundreds keV (5.4 σ C.L.)

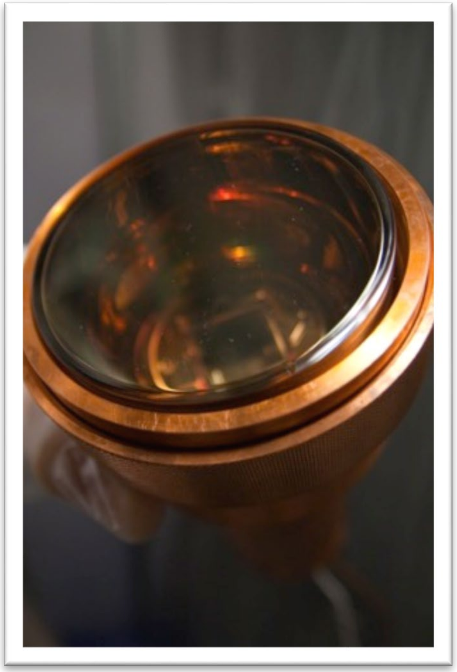
Further measurements with the same experimental set-up planned in the near future



Conclusions



- Model-independent evidence for a signal that satisfies all the requirement of the DM annual modulation signature at high C.L. (20 independent annual cycles with 3 different set-ups: 2.46 ton × yr)
- Modulation parameters determined with increasing precision
- New investigations on different peculiarities of the DM signal exploited in progress/foreseen



- Full sensitivity to many kinds of DM candidates and interactions types (both inducing recoils and/or e.m. radiation), **full sensitivity to low and high mass candidates**
- Model dependent analyses on new data allowed significantly improving the C.L. and restricting the allowed parameters' space for the various scenarios with respect to previous DAMA analysis
- DAMA/LIBRA–phase2 **continuing data taking**
- DAMA/LIBRA towards the lowering of the software energy threshold: **some efforts completed other are in progress**
- Continuing R&D on the development of **anisotropic scintillators** for the DM directionality approach
- Continuing investigations of **rare processes** other than DM