Inha University, Republic of Korea

Medium modification of Chiral Soliton Models

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Strategy and Motivation

How to construct a theoretical framework (model of ``nuclear physics")?

Our guiding principles are

- **simplicity (easy to analyse, transparent, etc…) <=> e.g. a small number terms in the Lagrangian;**
- **relation to phenomenology in an attractive way — as much as possible the peculiarities of strong interactions should be taken into account using as less as possible the number of parameters;**
- **universality <=> applicability to**
  - **hadron structure and spectrum studies (from light to heavy sector);**
  - **analysis of NN interactions;**
  - **nuclear many body problems <=> nucleonic systems (finite nuclei) and nuclear matter properties (EOS);**
  - **relation to mesonic atoms;**
  - **hadron structure changes in nuclear environment;**
  - **extreme density phenomena (e.g. neutron stars);**
  - **etc.**

Two possible ways:

- **to construct completely new approach;**
- **a bit fresh look to old ideas (e.g. putting a bit more phenomenological information).**

Studies

The studies were performed and going on in direction of

a single baryon properties

- in separate state considering it as a structure-full system
- nucleon in the community of their partners (EM and EMT form factors)
- nucleon in finite nuclei
- hyperons in nuclear matter (Thursday 4-B, 2:15 by Nam-Yong GHIM)
- heavy particles in nuclear matter (Thursday 4-B, 2:30 by Ho-Yeon WON)

as well as on the properties of the whole nucleonic systems

- infinite nuclear matter properties (volume and symmetry energy properties)
- matter under extreme conditions (e.g. neutron stars)
- matter with a strangeness
- few/many nucleon systems (symmetric nuclei, mirror nuclei, rare isotopes, halo nuclei,…)
- nucleon knock-out reactions (lepton-nucleus scattering)
- possible changes in in-medium NN interactions
- etc

Content

- Topological models
- Medium modifications
  - Mesons in nuclear matter
  - Baryons in nuclear matter
- Nuclear matter
- Neutron stars
- Non-spherically deformed nucleons
- Finite nuclei

Why topological models?

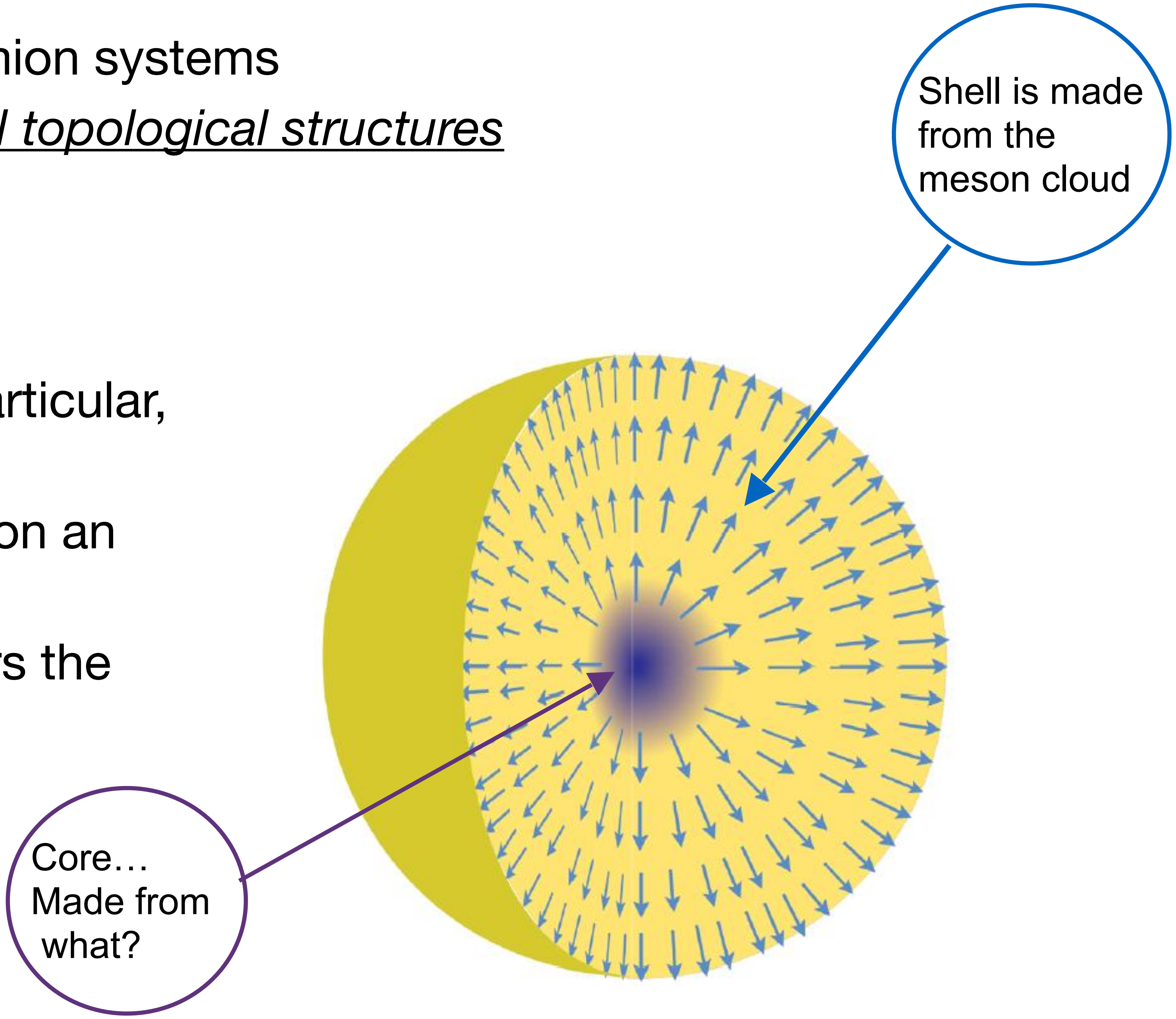
At fundamental level we may have

- fermions -> bosons are trivial fermion systems
- bosons -> fermions are *nontrivial topological structures*

Structure

From what made a nucleon and, in particular, its core?

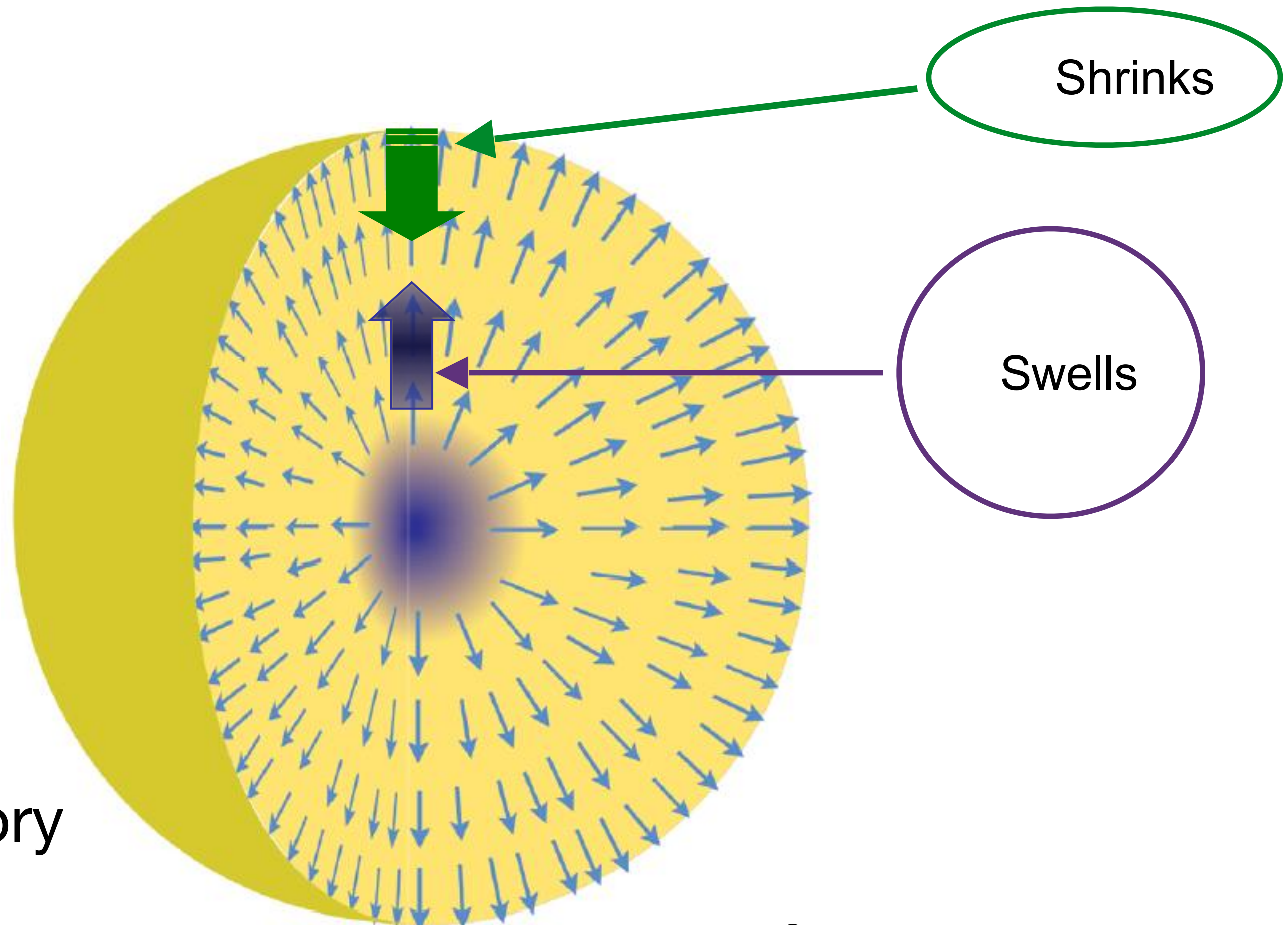
- The structure treatment depends on an energy scale
- At the limit of large number colours the core still has the mesonic content



Topological models

Stabilization mechanism

- Soliton has the finite size and the finite energy
- One needs at least two counter terms in the effective (mesonic) Lagrangian



Prototype: Skyrme model

[T.H.R. Skyrme, Pros.Roy.Soc.Lond. A260 (1961)]

- Nonlinear chiral effective meson (pionic) theory

$$\mathcal{L} = \frac{F_\pi^2}{16} \mathrm{Tr}\left(\partial_\mu U \partial^\mu U^\dagger\right) - \frac{1}{16e^2} \mathrm{Tr}\left[U^\dagger \partial_\mu U, U^\dagger \partial_\nu U\right]^2$$

Shrinking term　　　　Swelling term

- *Hedgehog* solution (nontrivial mapping)

$$U = \exp\left\{\frac{i\bar{\tau}\,\bar{\pi}}{2F_\pi}\right\} = \exp\{i\bar{\tau}\,\bar{n}F(r)\}$$

Topological models

The free space Lagrangian (was widely in use)
[G.S.Adkins et al. Nucl.Phys. B228 (1983)]

$$\mathcal{L} = \frac{F_\pi^2}{16}\,\mathrm{Tr}\left(\partial_\mu U \partial^\mu U^\dagger\right) - \frac{1}{16e^2}\,\mathrm{Tr}\left[U^\dagger\partial_\mu U, U^\dagger\partial_\nu U\right]^2 + \frac{F_\pi^2 m_\pi^2}{16}\,\mathrm{Tr}\left(U + U^\dagger - 2\right)$$

- Nontrivial structure: topologically stable solitons with the corresponding conserved topological number (baryon number) ***A***

$$U = \exp\{i\bar{\tau}\,\bar{\pi}/2F_\pi\} = \exp\{i\bar{\tau}\,\bar{n}F(r)\}$$

$$B^\mu = \frac{1}{24\pi^2}\varepsilon^{\mu\nu\alpha\beta}Tr(L_\nu L_\alpha L_\beta) \qquad L_\alpha = U^+\partial_\alpha U$$

$$A = \int d^3 r B^0$$

- Nucleon is quantized state of the classical soliton-skyrmion

$$H = M_{cl} + \frac{\bar{S}^2}{2I} = M_{cl} + \frac{\bar{T}^2}{2I},$$

$$|S = T, s, t\rangle = (-1)^{t+T}\sqrt{2T+1}D^{S=T}_{-t,s}(A)$$

Medium modifications

What happens in the nuclear medium?

The possible medium effects

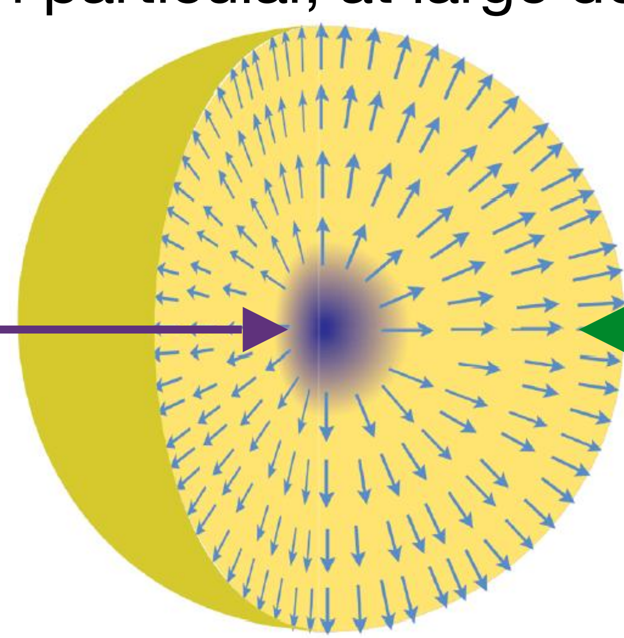
- Deformations (swelling or shrinking, multipole deformations) of nucleons
- Characteristic changes in: effective mass, charge distributions, all possible form factors
- NN interactions may change
- etc.

One should be able to describe all those phenomena

Soliton in the nuclear medium (phenomenological way)

- Outer shell modifications (informations from pionic atoms)
- Inner core modifications, in particular, at large densities (nuclear matter properties)

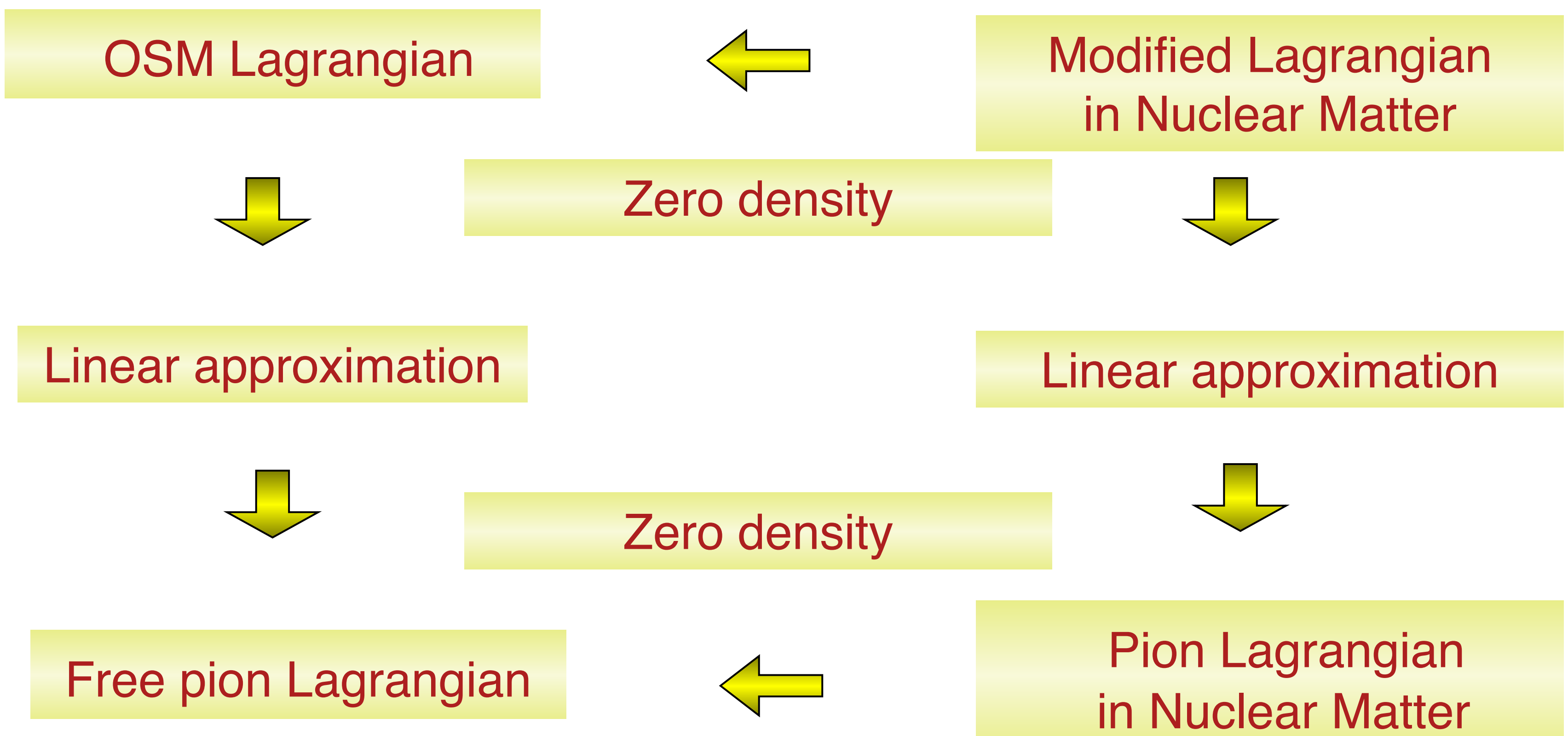
Inner core modifications in the nuclear medium may be related to:
- vector meson properties in the nuclear medium
- nuclear matter properties at saturation density



Meson cloud modifications in the nuclear medium:
Pion physics in the nuclear medium

Medium modifications

- Modifications of the mesonic sector modifies the baryonic sector
- Lagrangian satisfies some limiting conditions



Mesons in nuclear matter

"Outer shell" modifications

- In free space three types of pions can be treated separately: isospin breaking
- In nuclear matter: three types of polarization operators

$$(\partial^\mu \partial_\mu + m_\pi^2)\vec{\pi}^{(\pm,0)} = 0$$

$$(\partial^\mu \partial_\mu + m_\pi^2 + \hat{\Pi}^{(\pm,0)})\vec{\pi}^{(\pm,0)} = 0$$

$$\hat{\Pi}^0 = 2\omega U_{\rm opt} = \chi_s(\rho, b_0, c_0) + \vec{\nabla}\cdot\chi_p(\rho, b_0, c_0)\vec{\nabla}$$

$$\hat{\Pi}^0 = (\hat{\Pi}^- + \hat{\Pi}^+)/2, \qquad \hat{\Delta}\Pi = (\hat{\Pi}^- - \hat{\Pi}^+)/2$$

- Optic potential approach: parameters from the pion-nucleon scattering (including the isospin dependents)

	π-atom	$T_\pi = 50$ MeV
$b_0\ [m_\pi^{-1}]$	- 0.03	- 0.04
$b_1\ [m_\pi^{-1}]$	- 0.09	- 0.09
$c_0\ [m_\pi^{-3}]$	0.23	0.25
$c_1\ [m_\pi^{-3}]$	0.15	0.16
g'	0.47	0.47

Medium modifications

"Outer shell" modifications in the Lagrangian [U.Meissner *et al.*, EPJ A36 (2008)]

$$\mathcal{L}_2^* = \frac{F_\pi^2}{16}\alpha_\tau \mathrm{Tr}\left(\partial_0 U \partial_0 U^\dagger\right) - \frac{F_\pi^2}{16}\alpha_s \mathrm{Tr}\left(\partial_i U \partial_i U^\dagger\right)$$

$$\mathcal{L}_m^* = -\frac{F_\pi^2 m_\pi^2}{16}\alpha_m \mathrm{Tr}\left(2 - U - U^\dagger\right)$$

- Due to the nonlocality of optic potential the kinetic term is also modified
- Due to energy and momentum dependence of the optic potential parameters the following parts of the kinetic term are modified in different forms:
  - Temporal part
  - Space part

$$\hat{\Pi}^0 = 2\omega U_{\mathrm{opt}} = \chi_s(\rho, b_0, c_0) + \vec{\nabla}\cdot\chi_p(\rho, b_0, c_0)\vec{\nabla}$$

	π-atom	$T_\pi = 50$ MeV
b_0 $[m_\pi^{-1}]$	- 0.03	- 0.04
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g'	0.47	0.47

“Inner core” modifications
[UY & H.Ch. Kim, PRC83 (2011); UY, PRC88 (2013)]

$$\mathcal{L}_4^* = -\frac{1}{16e^2\zeta_\tau}\mathrm{Tr}\left[U^\dagger\partial_0 U, U^\dagger\partial_i U\right]^2 + \frac{1}{32e^2\zeta_s}\mathrm{Tr}\left[U^\dagger\partial_i U, U^\dagger\partial_j U\right]^2$$

may be related to

- Vector meson properties in nuclear matter
- Nuclear matter properties

$$\zeta_{\tau,s} = \zeta_{\tau,s}(\rho, \delta\rho, \mathrm{parameters})$$

Medium modifications

Final Lagrangian
[UY, JKPS62 (2013); UY, PRC88 (2013)]

Separated into two parts

$$\mathcal{L}^* = \mathcal{L}^*_{\rm sym} + \mathcal{L}^*_{\rm asym}$$

- Isoscalar part

$$\mathcal{L}^*_{\rm sym} = \mathcal{L}^*_2 + \mathcal{L}^*_4 + \mathcal{L}^*_m$$

- Isovector part

$$\mathcal{L}^*_{\rm asym} = \mathcal{L}^*_{\delta m} + \mathcal{L}^*_{\delta\rho}$$

- **Nuclear matter stabilization** (→ $\mathcal{L}^*_2$, $\mathcal{L}^*_4$)
- **Asymmetric matter properties** (→ $\mathcal{L}^*_{\delta\rho}$)

$$\mathcal{L}^*_2 = \frac{F_\pi^2}{16}\,\alpha_\tau {\rm Tr}\left(\partial_0 U \partial_0 U^\dagger\right) - \frac{F_\pi^2}{16}\,\alpha_s {\rm Tr}\left(\partial_i U \partial_i U^\dagger\right)$$

$$\mathcal{L}^*_4 = -\frac{1}{16e^2\zeta_\tau}\,{\rm Tr}\left[U^\dagger\partial_0 U, U^\dagger\partial_i U\right]^2 + \frac{1}{32e^2\zeta_s}\,{\rm Tr}\left[U^\dagger\partial_i U, U^\dagger\partial_j U\right]^2$$

$$\mathcal{L}^*_m = -\frac{F_\pi^2 m_\pi^2}{16}\,\alpha_m {\rm Tr}\left(2 - U - U^+\right)$$

$$\mathcal{L}^*_{\delta m} = -\frac{F_\pi^2}{32}\sum_{a=1}^{2}(m^2_{\pi^\pm} - m^2_{\pi^0}){\rm Tr}\left(\tau_a U\right){\rm Tr}\left(\tau_a U^\dagger\right)$$

$$\mathcal{L}^*_{\delta\rho} = -\frac{F_\pi^2}{16}\,m_\pi\alpha_e\ \varepsilon_{ab3}{\rm Tr}\left(\tau_a U\right){\rm Tr}\left(\tau_b \partial_0 U^\dagger\right)$$

Medium modifications

Reparametrization
[UY, PRC88 (2013)]

- Five density dependent parameters

- Rearrangment (technical simplification to describe nuclear matter)

Shell modifications

$$F_{\pi,\tau} \to F^*_{\pi,\tau}, \quad F_{\pi,s} \to F^*_{\pi,s},$$

Core modifications

$$e_\tau \to e^*_\tau, \quad e_s \to e^*_s$$

$$m_\pi \to m^*_\pi,$$

$$+C_1 \frac{\rho}{\rho_0} = f_1\left(\frac{\rho}{\rho_0}\right) \equiv \sqrt{\frac{\alpha_p^0}{\gamma_s}}$$

$$+C_2 \frac{\rho}{\rho_0} = f_2\left(\frac{\rho}{\rho_0}\right) \equiv \frac{\alpha_s^{00}}{(\alpha_p^0)^2 \gamma_s}$$

$$+C_3 \frac{\rho}{\rho_0} = f_3\left(\frac{\rho}{\rho_0}\right) \equiv \frac{(\alpha_p^0 \gamma_s)^{3/2}}{\alpha_s^{02}}$$

$$\frac{\alpha_e}{\gamma_s} = f_4\left(\frac{\rho}{\rho_0}\right)\frac{\rho_n - \rho_p}{\rho_0} = \frac{C_4 \frac{\rho}{\rho_0}}{1 + C_5 \frac{\rho}{\rho_0}} \frac{\rho_n - \rho_p}{\rho_0}$$

Nucleon in nuclear matter

Structure studies 1: Energy momentum tensor

- It allows to address the questions like:
  - How are the total angular momentum and angular momentum of the nucleon shared among its constituents?
  - How are the strong forces experienced by its constituents distributed inside the nucleon?
- EMT form factors studied in lattice QCD, ChPT and in different models (chiral quark soliton model, Skyrme model, etc.)
- We made further step studying EMT form factors in nuclear matter

Structure studies 1: Energy momentum tensor

- Definition

$$\langle p'|\hat{T}_{\mu\nu}(0)|p\rangle = \bar{u}(p',\, s')\left[M_2(t)\,\frac{P_\mu P_\nu}{M_N} + J(t)\,\frac{i(P_\mu \sigma_{\nu\rho} + P_\nu \sigma_{\mu\rho})\Delta^\rho}{2M_N} + d_1(t)\,\frac{\Delta_\mu \Delta_\nu - g_{\mu\nu}\Delta^2}{5M_N}\right] u(p,\, s)\,,$$

- Three form factors give an information about energy distribution, angular momentum distribution and about the stabilization of strong forces inside the nucleon

$$T^*_{00}(r) = \frac{F^{*2}_{\pi,s}}{8}\left(\frac{2\sin^2 F}{r^2} + F'^2\right) + \frac{\sin^2 F}{2\,e^{*2}\,r^2}\left(\frac{\sin^2 F}{r^2} + 2F'^2\right) + \frac{m^{*2}_\pi F^{*2}_{\pi,s}}{4}\,(1-\cos F)\,,$$

$$T^*_{0k}(\boldsymbol{r},\boldsymbol{s}) = \frac{\epsilon^{klm} r^l s^m}{(\mathrm{s}\times\mathrm{r})^2}\,\rho^*_J(r)\,,$$

$$T^*_{ij}(\boldsymbol{r}) = s^*(r)\left(\frac{r_i r_j}{r^2} - \frac{1}{3}\,\delta_{ij}\right) + p^*(r)\,\delta_{ij}$$

$$M^*_2(t) - \frac{t}{5M^{*2}_N}\,d^*_1(t) = \frac{1}{M^*_N}\int \mathrm{d}^3\boldsymbol{r}\; T^*_{00}(r)\; j_0(r\sqrt{-t})\,,$$

$$d^*_1(t) = \frac{15M^*_N}{2}\int \mathrm{d}^3\boldsymbol{r}\; p^*(r)\;\frac{j_0(r\sqrt{-t})}{t}\,,$$

$$M^*_2(0) = \frac{1}{M^*_N}\int \mathrm{d}^3\boldsymbol{r}\; T^*_{00}(r) = 1\,, \quad J^*(0) = \int \mathrm{d}^3\boldsymbol{r}\; \rho^*_J(r) = \frac{1}{2}\,.$$

$$J^*(t) = 3\int \mathrm{d}^3\boldsymbol{r}\; \rho^*_J(r)\;\frac{j_1(r\sqrt{-t})}{r\sqrt{-t}}\,,$$

Structure studies1: Energy momentum tensor related quantities
[H.C.Kim, P. Schweitzer, UY, Phys. Lett. B718 (2012)]

Different quantities related to the nucleon EMT densities and their form factors: $T^*_{00}(0)$ denotes the energy in the center of the nucleon; $\langle r^2_{00}\rangle^*$ and $\langle r^2_J\rangle^*$ depict the mean square radii for the energy and angular momentum densities, respectively; $p^*(0)$ represents the pressure in the center of the nucleon, whereas r^*_0 designates the position where the pressure changes its sign; d^*_1 is the value of the $d^*_1(t)$ form factor at the zero momentum transfer.

ρ/ρ_0	$T^*_{00}(0)$ [GeV fm^{-3}]	$\langle r^2_{00}\rangle^*$ [fm^2]	$\langle r^2_J\rangle^*$ [fm^2]	$p^*(0)$ [GeV fm^{-3}]	r^*_0 [fm]	d^*_1
0	1.45	0.68	1.09	0.26	0.71	−3.54
0.5	0.96	0.83	1.23	0.18	0.82	−4.30
1.0	0.71	0.95	1.35	0.13	0.90	−4.85

Structure studies 1: Pressure distribution inside the nucleon in free space and in symmetric matter [H.C.Kim, P. Schweitzer, UY, Phys. Lett. B718 (2012)]

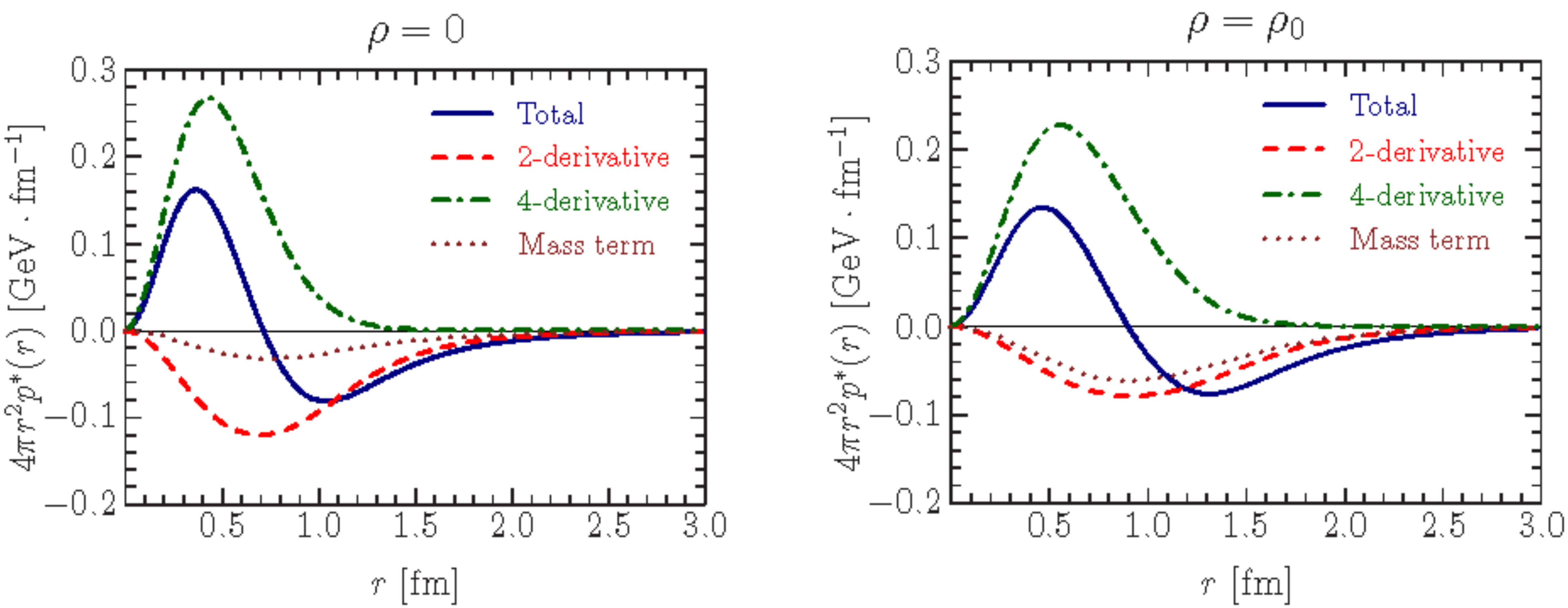


FIG. 3: (Color online) The decomposition of the pressure densities $4\pi r^2 p^*(r)$ as functions of r, in free space ($\rho = 0$) and at $\rho = \rho_0$, in the left and right panels, respectively. The solid curves denote the total pressure densities, the dashed ones represent the contributions of the 2-derivative (kinetic) term, the long-dashed ones are those of the 4-derivative (stabilizing) term, and the dotted ones stand for those of the pion mass term.

Nucleon in nuclear matter

Structure studies 2: Transverse EM charge densities

- Definition of EM FF's $\langle N(p', S') | J_\mu^{EM}(0) | N(p, S) \rangle$

$$= \bar{u}_N(p', S') \left[\gamma_\mu F_1^*(q^2) + i \frac{\sigma_{\mu\nu} q^\nu}{2m_N} F_2^*(q^2) \right] u_N(p, S).$$

- These Pauli and Dirac FF's can be expressed by Sachs FF's

$$G_E(Q^2) = F_1(Q^2) + \frac{Q^2}{4M_N^2} F_2(Q^2)$$
$$G_M(Q^2) = F_1(Q^2) + F_2(Q^2),$$

- They give an information about transverse charge distributions inside the nucleon

$$\rho_0^*(b) = \int_0^\infty \frac{Q\,dQ}{2\pi} J_0(bQ) \frac{G_E^*(Q^2) + \tau G_M^*(Q^2)}{1+\tau}$$

$$\rho_T^*(\boldsymbol{b}) = \rho_0^*(b) - \sin(\phi_b - \phi_S) \times \int_0^\infty \frac{Q^2\,dQ}{4\pi m_N} J_1(bQ) \frac{-G_E^*(Q^2) + G_M^*(Q^2)}{1+\tau},$$

$$\boldsymbol{b} = b(\cos\phi_b \hat{\boldsymbol{e}}_x + \sin\phi_b \hat{\boldsymbol{e}}_y)$$

Nucleon in nuclear matter

Structure studies 2: Transverse EM charge densities inside an unpolarized nucleon [UY, H.C.Kim, Phys.Lett. B726 (2013)]

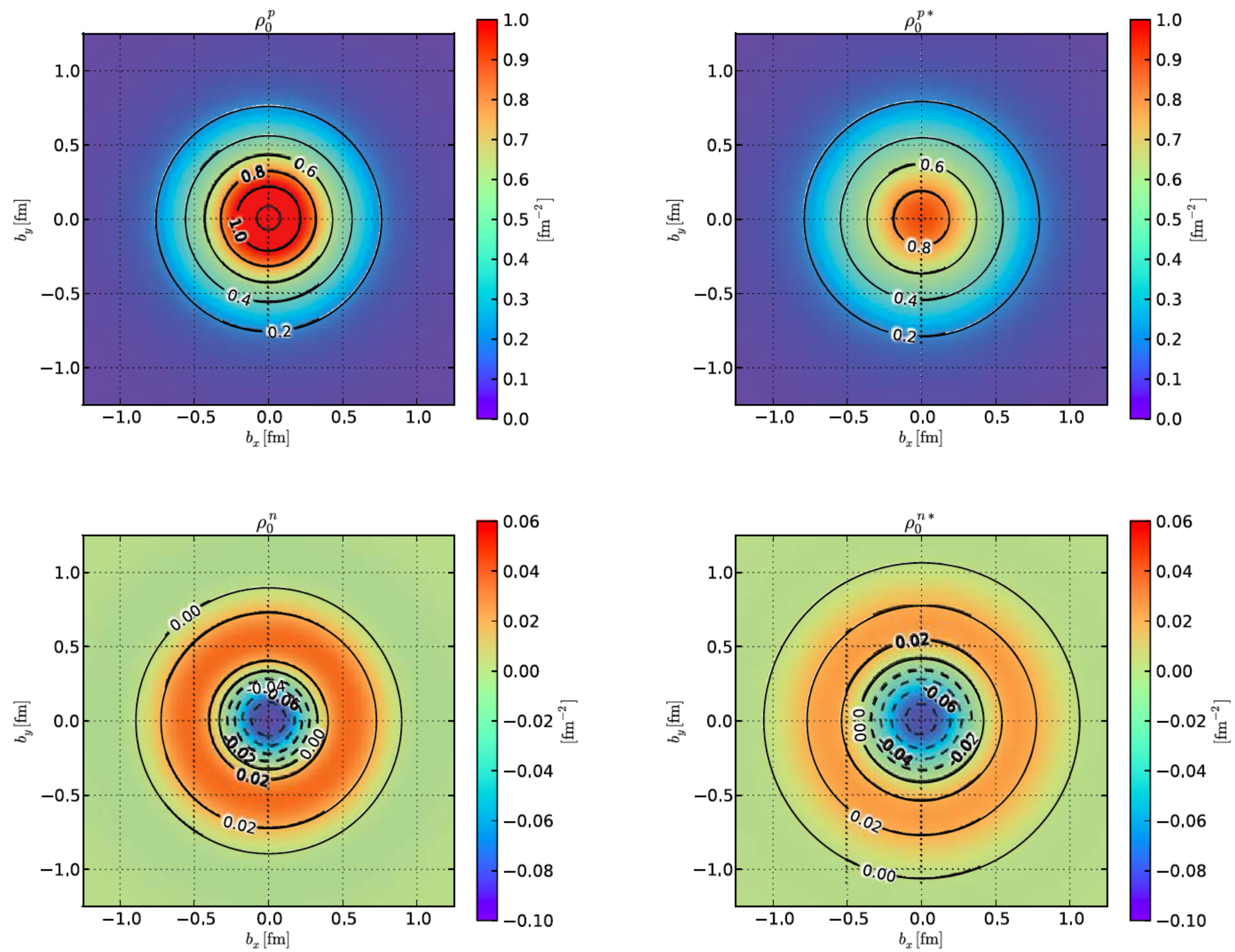


Fig. 3. Quark transverse charge densities inside an unpolarized proton (upper panels) and a neutron (lower panels) in free space (left panels) and at nuclear matter density $\rho_0 = 0.5m_\pi^3$ (right panels).

Structure studies 2: Transverse EM charge densities inside the polarized nucleon [UY, H.C.Kim, Phys. Lett. B726 (2013)]

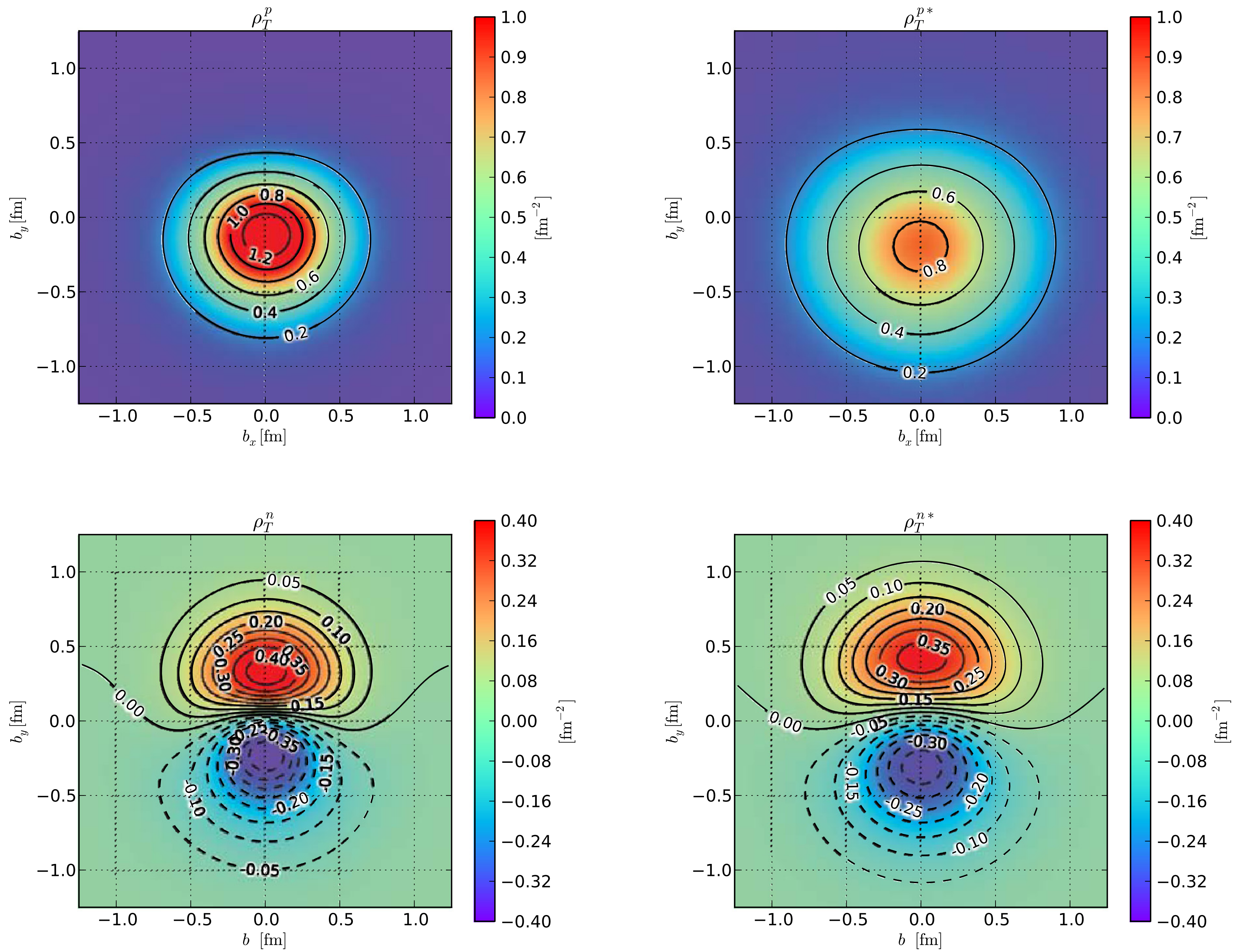


Fig. 4. Transverse charge densities of the proton (upper panels) and neutron (lower panels) in free space (left panels) and in nuclear matter with the density $\rho_0 = 0.5 m_\pi^3$ (right panels).

Nucleon in nuclear matter

Masses [UY, PRC88 (2013)]

- Isoscalar effective mass

$$m^*_{N,s} = M^*_S + \frac{3}{8\Lambda^*} + \frac{\Lambda^*}{2}\left(a^{*2} + \frac{\Lambda^{*2}_{\text{env}}}{\Lambda^{*2}}\right)$$

- Isovector effective mass (relevant to: universe evolution in Early stage; stability of drip line nuclei; mirror nuclei; transport in heavy-ion collisions; asymmetric nuclear matter)

$$\Delta m^*_{np} = a^* + \frac{\Lambda^*_{\text{env}}}{\Lambda^*}$$

- Effective masses of the nucleons

$$m^*_{n,p} = m^*_{N,s} - \Delta m^*_{np} T_3$$

Nuclear matter

From the Bethe-Weizsacker formula

$$\varepsilon\,(A,Z) = -a_V + a_S \frac{(N-Z)^2}{A^2} +$$ [illegible]

The binding-energy-formula terms in the framework of present model can be obtained considering

We reproduced (Volume term, Asymmetry term)

- Volume term
  - Symmetric infinite nuclear matter
- Asymmetry term
  - Isospin asymmetric environment
- Surface and Coulomb terms
  - Nucleons in a finite volume
- Finite nuclei properties
  - Local density approximation

The volume term and Symmetry energy

- At infinite nuclear matter approximation the binding energy per nucleon takes the form

$$\varepsilon(\lambda,\delta) = \varepsilon_V(\lambda) + \varepsilon_S\delta^2 + O(\delta^4) \equiv \varepsilon_V(\lambda) + \varepsilon_A(\lambda,\delta)$$

  - λ is normalised nuclear matter density
  - δ is asymmetry parameter
  - ε_S is symmetry energy

- In our model
  - Symmetric matter
  - Asymmetric matter

$$\varepsilon_V(\lambda) = m^*_{N,s}(\lambda,0) - m^{\text{free}}_N$$

$$\varepsilon_A(\lambda,\delta) = \varepsilon(\lambda,\delta) - \varepsilon_V(\lambda)$$
$$= m^*_{N,s}(\lambda,\delta) - m^*_{N,s}(\lambda,0) + m^*_{N,V}(\lambda,\delta)\delta$$

Nuclear matter properties

- Symmetric matter properties (pressure, compressibility and third derivative)

$$p = \rho_0 \lambda^2 \left. \frac{\partial \varepsilon_V(\lambda)}{\partial \lambda} \right|_{\lambda=1}, \quad K_0 = 9\rho^2 \left. \frac{\partial^2 \varepsilon_V(\lambda)}{\partial \rho^2} \right|_{\rho=\rho_0} \quad Q = 27\lambda^3 \left. \frac{\partial^3 \varepsilon_V(\lambda)}{\partial \lambda^3} \right|_{\lambda=1}$$

- Symmetry energy properties (coefficient, slop and curvature)

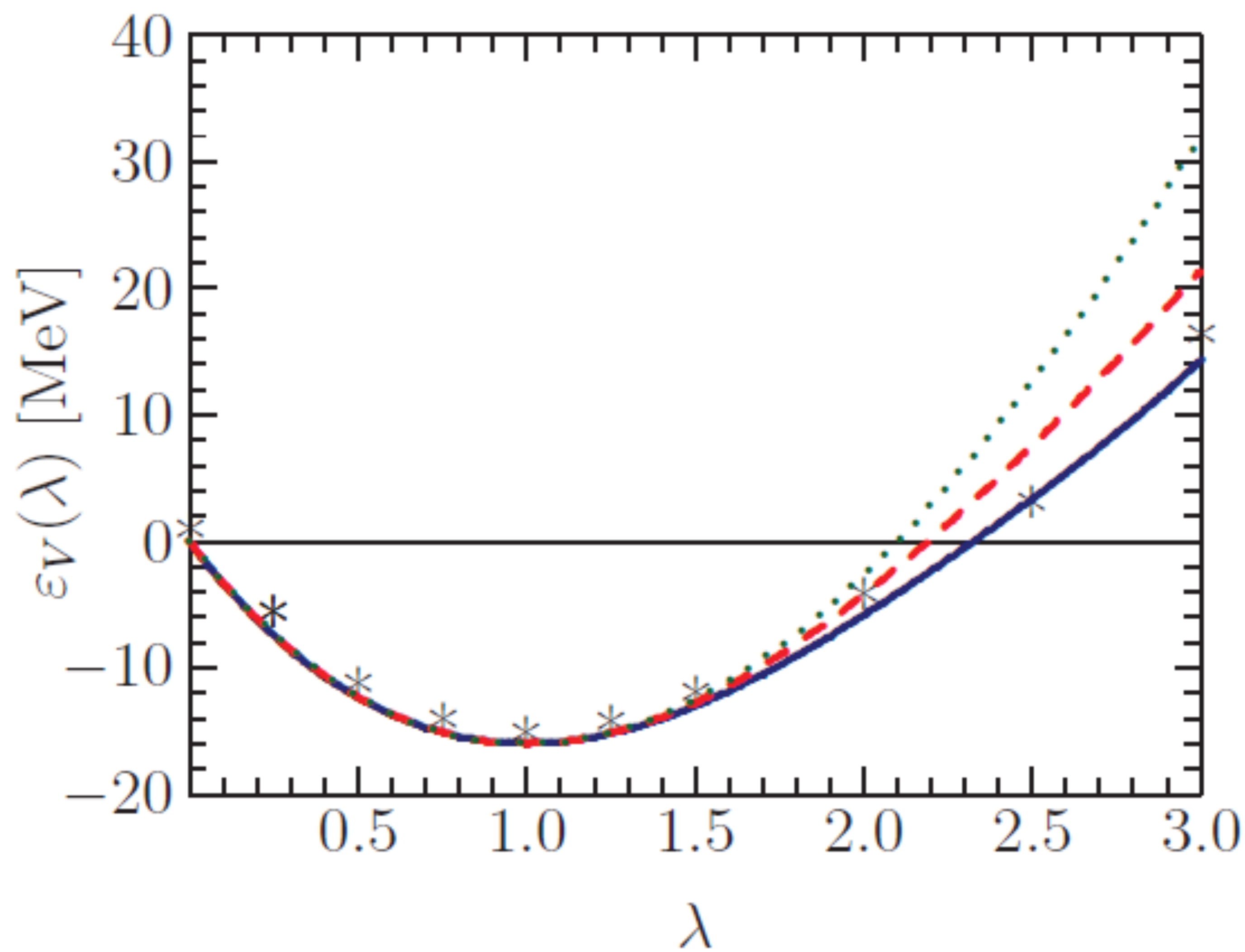
$$\varepsilon_s(\lambda) = \varepsilon_s(1) + \frac{L_s}{3}(\lambda - 1) + \frac{K_S}{18}(\lambda - 1)^2 + \cdots$$

Symmetric matter

Volume energy
[UY, PRC88 (2013)]

- Set I – solid
- Set II – dashed
- Set III – dotted

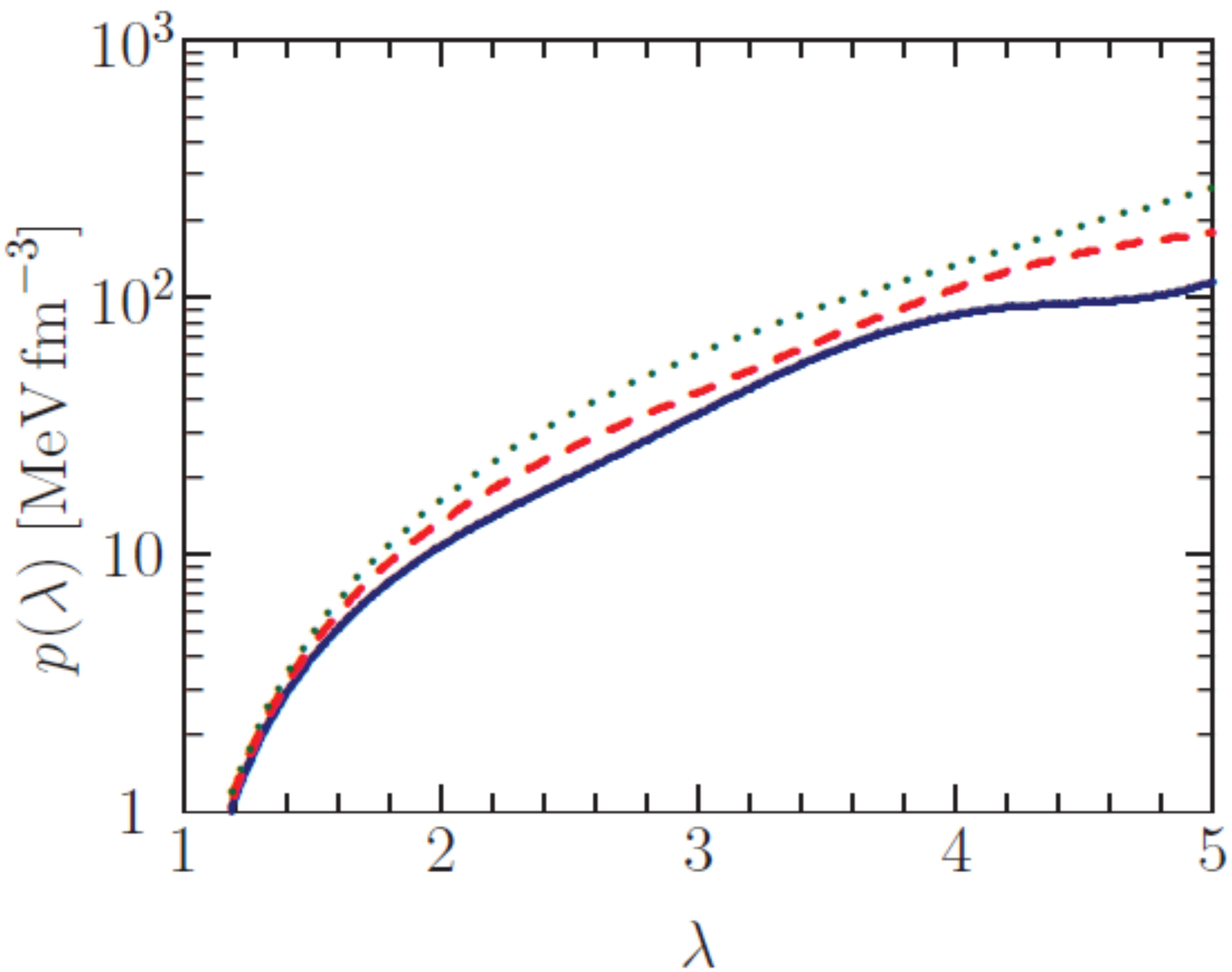
For comparison: Akmal-Pandharipande-Ravenhall (APR) predictions
[PRC 58, 1804 (1998)] are given by stars.
(From Arigonna 2 body interactions + 3 body interactions)

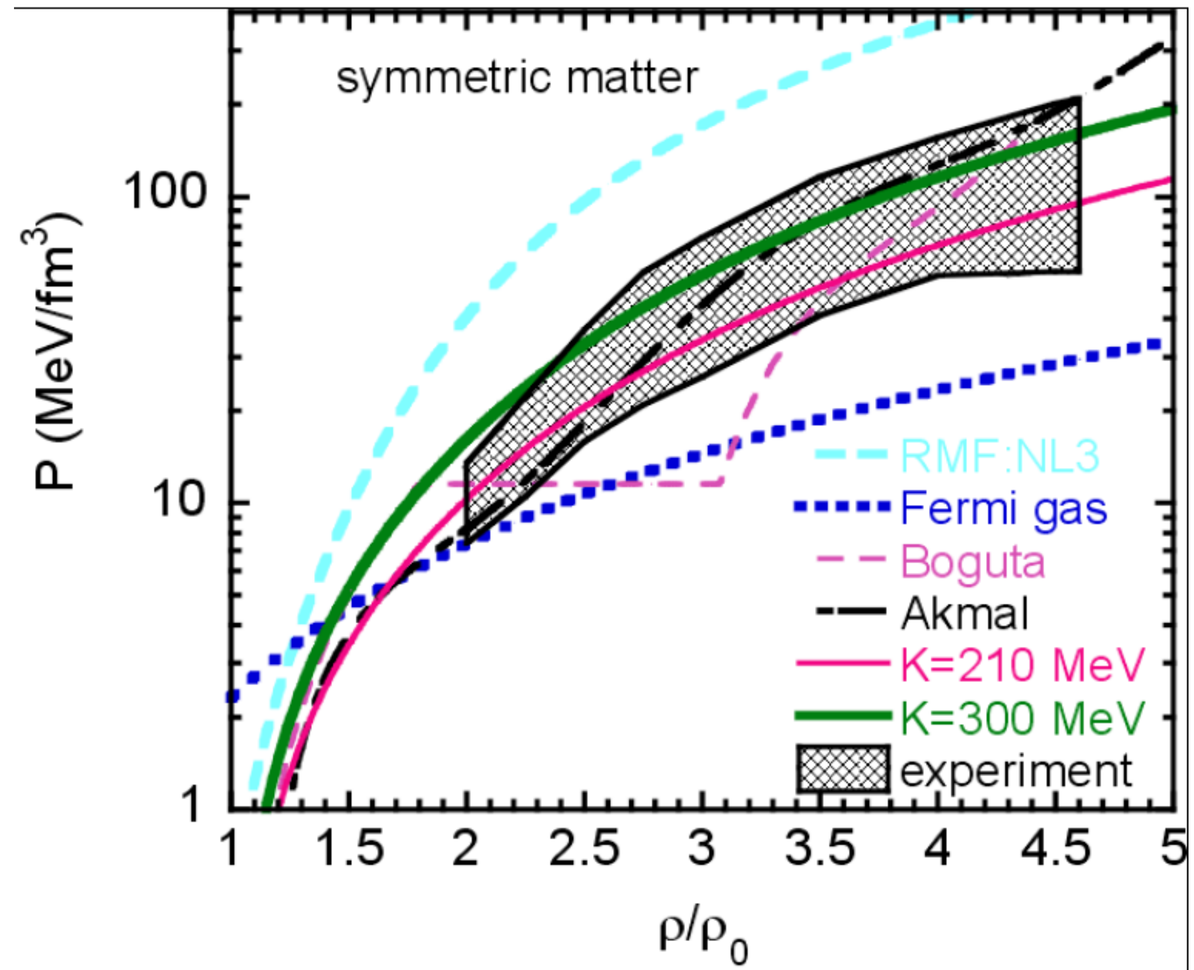


Set	C_1	C_2	C_3	$\varepsilon_V(\rho_0)$ (MeV)	K_0 (MeV)	Q (MeV)
I	−0.279	0.737	1.782	−16	240	−410
II	−0.273	0.643	1.858	−16	250	−279
III	−0.277	0.486	2.124	−16	260	−178

Symmetric matter

Pressure
[UY, PRC88 (2013)]





For comparison: Right figure from
Danielewicz- Lacey-Lynch, Science 298, 1592 (2002).
(Deduced from experimental flow data and simulations studies)

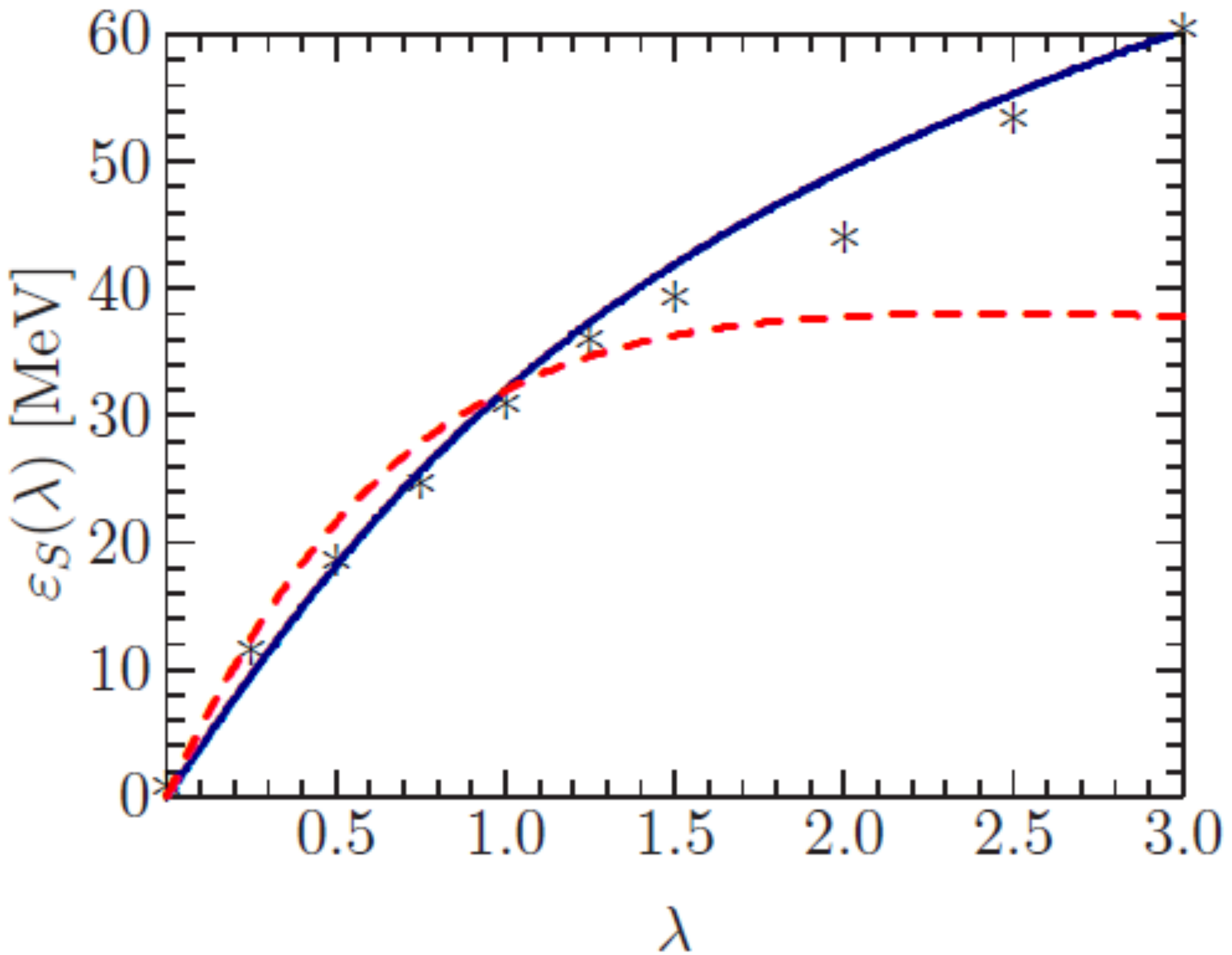
Asymmetric matter

Symmetry energy

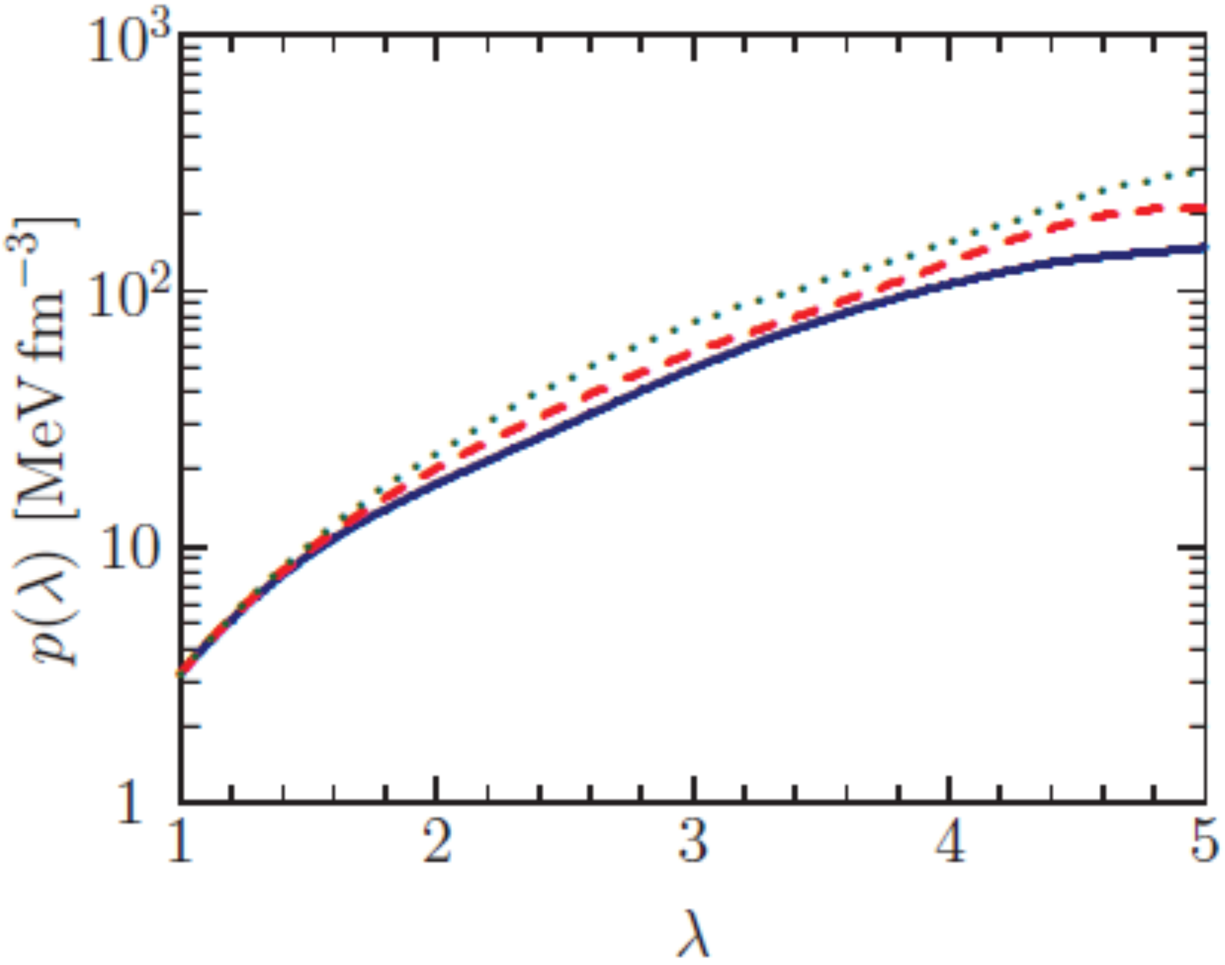
- Solid $L_s = 70\,\text{MeV}$
- Dashed $L_s = 40\,\text{MeV}$

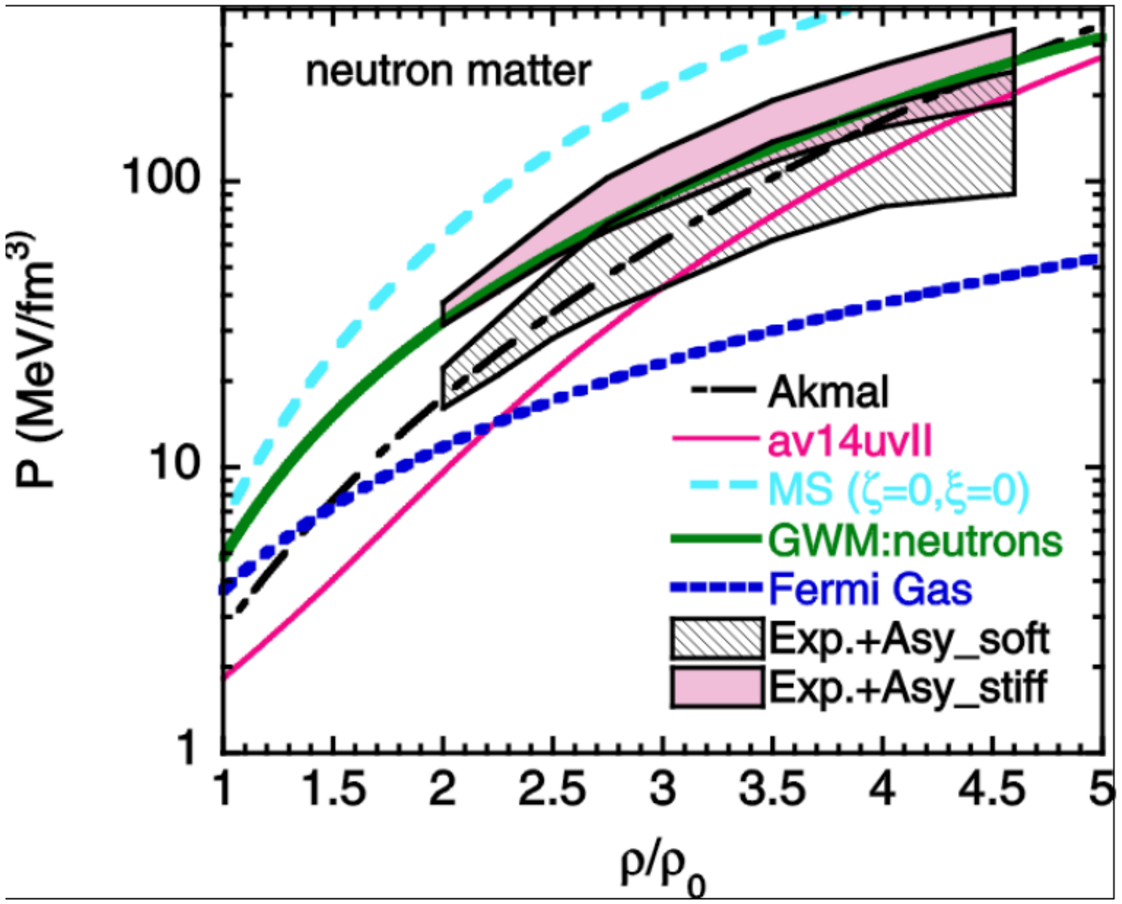
For comparison: Akmal-Pandharipande-Ravenhall (APR) predictions [PRC 58, 1804 (1998)] are given by stars.

(From arigonna 2 body interactions + 3 body interactions)



Pressure in neutron matter
[UY, PRC88 (2013)]





For comparison: Right figure from
Danielewicz- Lacey-Lynch, Science 298, 1592 (2002).
(Deduced from experimental flow data and simulations studies)

Low density behaviour of symmetry energy

For comparison:
Trippa-Colo-Vigezzi
[PRC 77, 061304 (2008)];
From analysis of GDR
(208Pb).

$$23.3 < \varepsilon_s(\rho = 0.1\mathrm{fm}^{-3}) < 24.9\ \mathrm{MeV}$$

Consequently one can predict in this model:

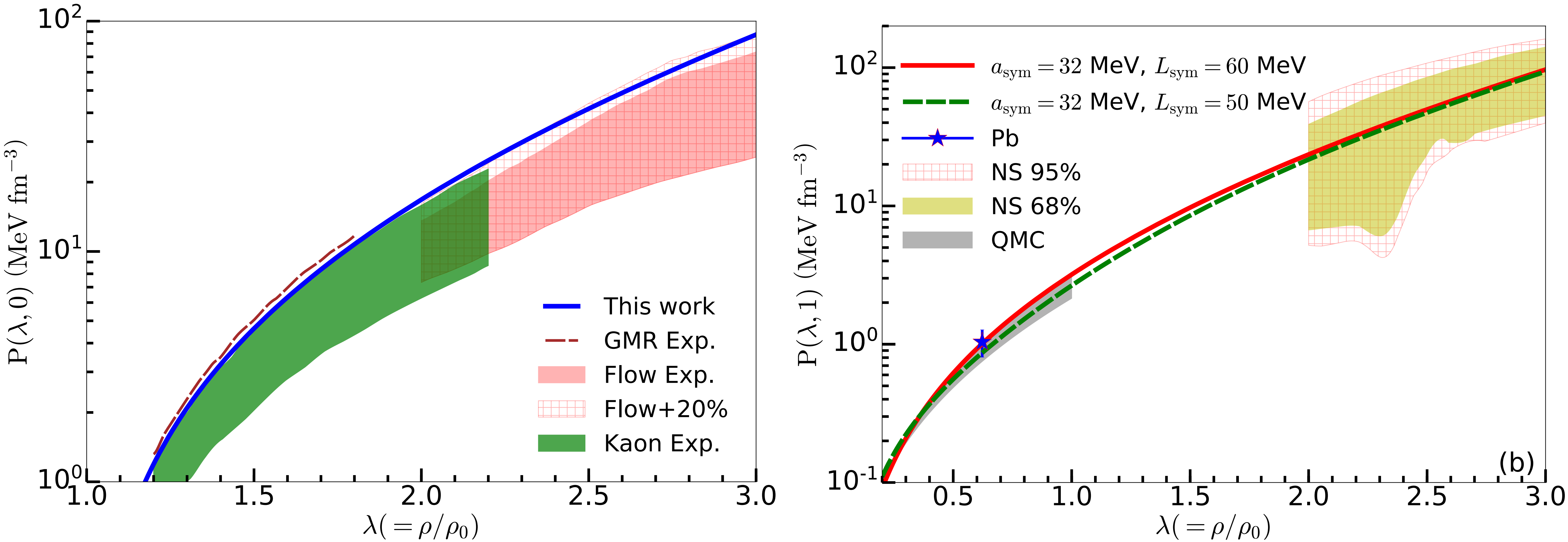
$$K_\tau = K_s - 6L_s$$

$$K_{0,2} = K_\tau - \frac{Q}{K_0} L_S$$

$\varepsilon_S(\rho_0)$ [MeV]	L_S [MeV]	K_S [MeV]	K_τ [MeV]	$K_{0,2}$ [MeV]	$\varepsilon_S(0.1\mathrm{fm}^{-3})$ [MeV]
32	40	−181	−301	−257	25.15
32	50	−160	−310	−254	24.15
32	60	−126	−306	−239	23.22
32	70	−80	−290	−211	22.37
32	80	−21	−261	−172	21.57
32	90	50	−220	−119	20.82
32	100	134	−166	−55	20.13

Nuclear matter (SU(3) model independent approach with hyperons) Thursday 4-B, 2:15 by Nam-Yong GHIM

Pressure [N.Y.Ghim, G.S.Yang, H.Ch.Kim, UY, PRC103 (2021)]



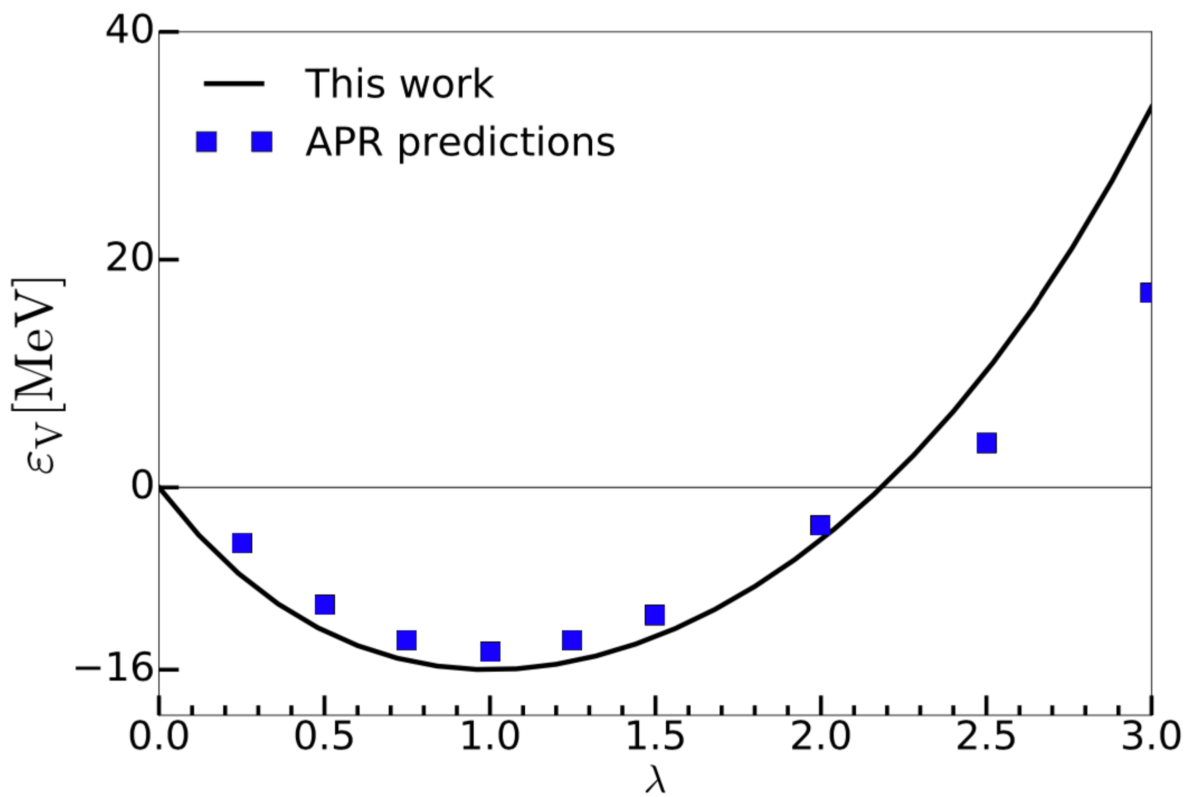
[6] W. G. Lynch, M. B. Tsang, Y. Zhang, P. Danielewicz, M. Famiano, Z. Li and A. W. Steiner, Prog. Part. Nucl. Phys. **62**, 427 (2009).

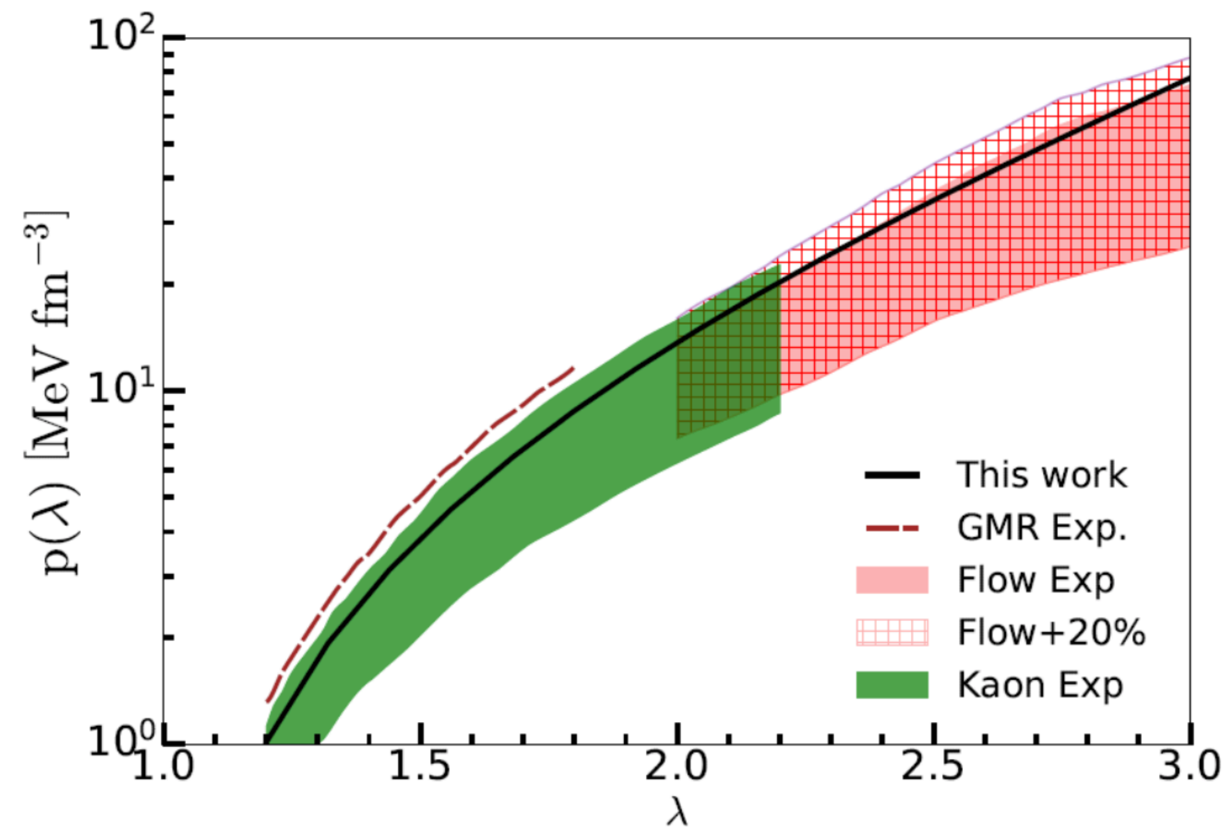
[7] A. W. Steiner, J. M. Lattimer and E. F. Brown, Astrophys. J. Lett. **765**, L5 (2013).

[8] M. B. Tsang et al., Phys. Rev. C 86, 015803 (2012).

Nuclear matter (SU(3) model with heavy baryons)
Thursday 4-B, 2:30 by Ho-Yeon WON

Volume energy and Pressure [H.Y.Won, UY, H.Ch.Kim, 2110.04561 [nucl-th]]





Nuclear matter (SU(3) model with heavy baryons)
Thursday 4-B, 2:30 by Ho-Yeon WON

Charmed baryons in nuclear matter [H.Y.Won, UY, H.Ch.Kim, 2110.04561 [nucl-th]]

$$M^* \equiv M_D^* = (1 + C_4\lambda)M_D \qquad \tau_{\text{heavy}}^* = (1 + C_4\lambda)\tau_{\text{heavy}}$$

Baryon	$M_B,\ \rho = 0$		$\Delta M_B,\ \rho = \rho_0$		
	Exp. [5]	This work	$C_4 = -0.1$	$C_4 = 0$	$C_4 = 0.1$
Λ_c	2286.5	2286.0	-166.91	21.22	209.34
Ξ_c	2469.4	2437.8	-132.30	54.48	241.25
Σ_c	2453.5	2564.5	-86.50	101.54	289.57
Ξ_c'	2576.8	2646.8	-69.89	117.27	304.43
Ω_c	2695.2	2721.6	-54.23	132.17	318.56

$$\Delta M_B = M_B^* - M_B$$

Neutron stars

Neutron star properties

- TOV equations

$$-\frac{dP(r)}{dr} = \frac{G\mathcal{E}(r)\mathcal{M}(r)}{r^2}\left(1-\frac{2G\mathcal{M}(r)}{r}\right)^{-1}\left(1+\frac{P(r)}{\mathcal{E}(r)}\right)\left(1+\frac{4\pi r^3 P(r)}{\mathcal{M}(r)}\right)$$

- Energy-pressure relation

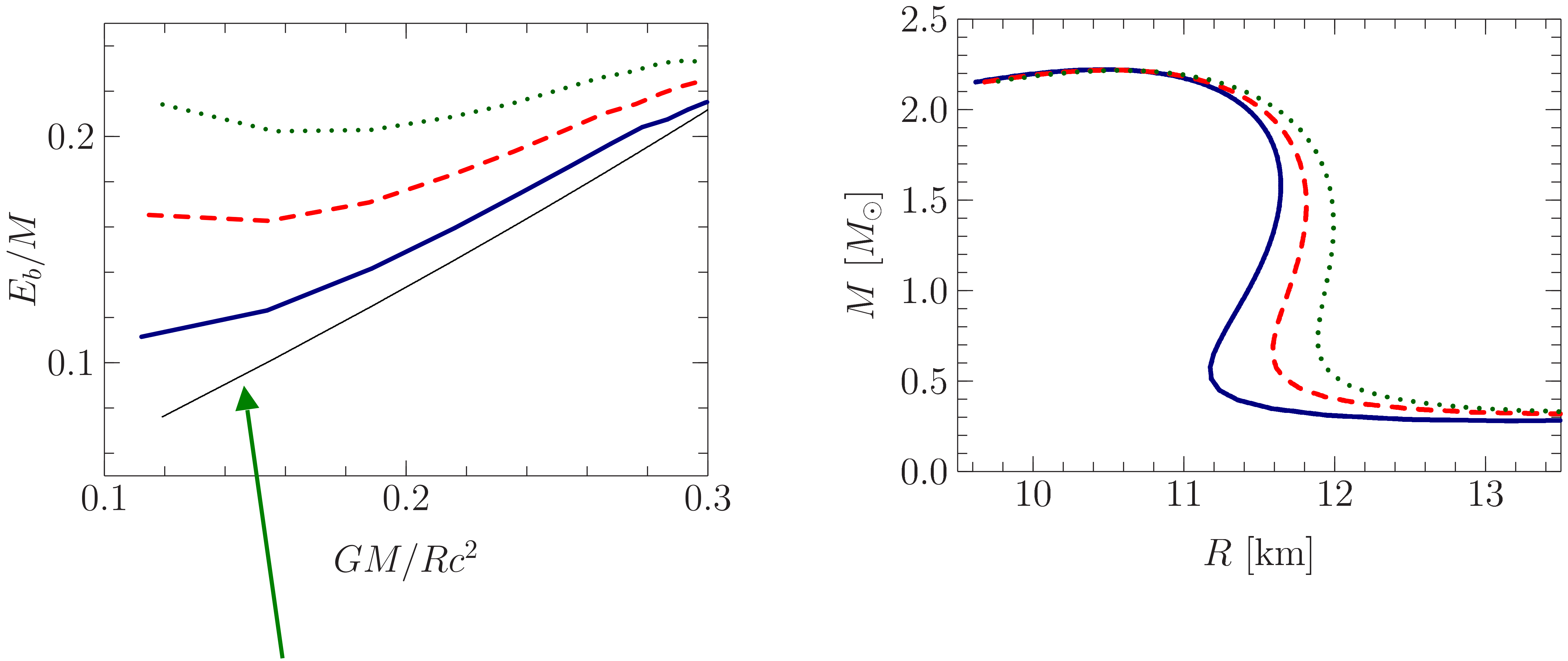
$$P = P(\mathcal{E})$$

$$P(\lambda) = \rho_0\lambda^2\frac{\partial\varepsilon(\lambda,1)}{\partial\lambda}\,,$$
$$\mathcal{E}(\lambda) = [\varepsilon(\lambda,1)+m_N]\lambda\rho_0.$$

- Neutron star's mass

$$\mathcal{M}(r) = 4\pi\int_0^r \mathrm{d}r\, r^2\mathcal{E}(r)\,.$$

Neutron star properties [UY, PLB749 (2015)]



From Ref. [J.M. Lattimer & M. Prakash, Astrophys. J. 550 (2001)].

Neutron stars

Neutron star properties [UY, PLB749 (2015)]

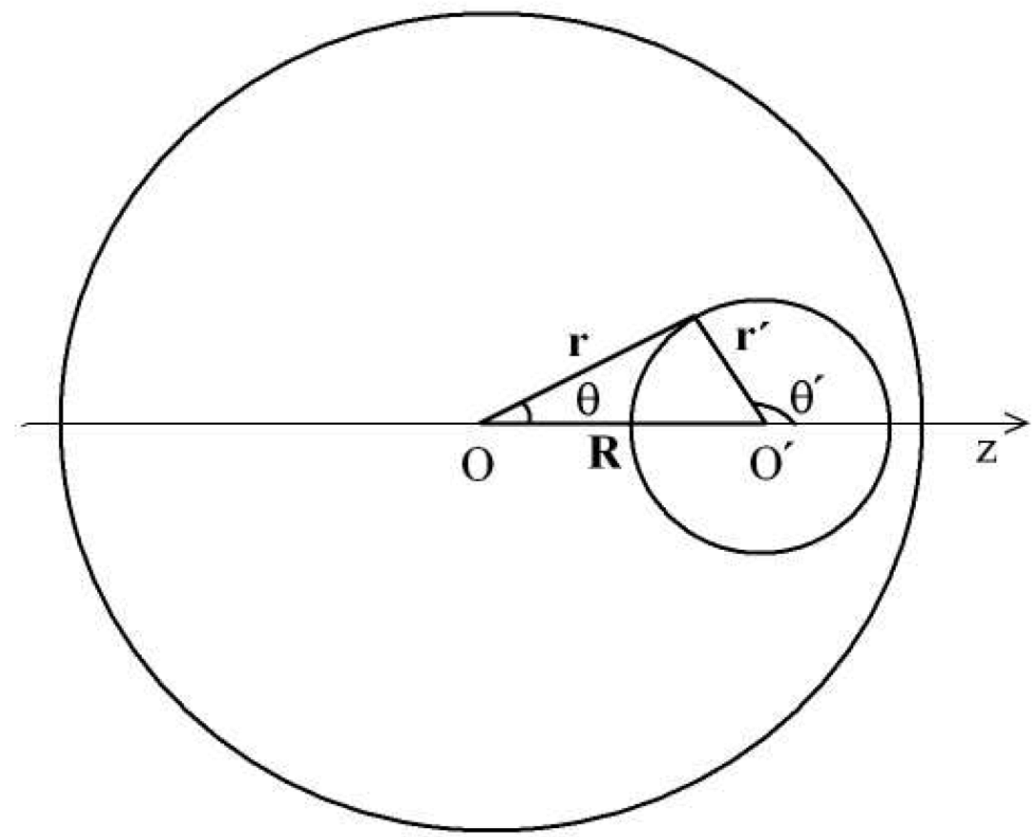
TABLE III: Properties of the neutron stars from the different sets of parameters (see Tables I and II for the values of parameters): n_c is central number density, ρ_c is central energy-mass density, R is radius of the neutron star, $M_{\max}$ is possible maximal mass, A is number of baryons in the star, E_b is binding energy of the star. In the left panel we represent the neutron star properties corresponding to the maximal mass $M_{\max}$ and in right panel approximately 1.4 solar mass neutron star properties. The last two lines are results from the Ref. [21].

Set	n_c [fm^{-3}]	ρ_c [10^{15}gr/cm^3]	R [km]	$M_{\max}$ [$M_\odot$]	A [10^{57}]	E_b [10^{53}erg]	n_c [fm^{-3}]	ρ_c [10^{15}gr/cm^3]	R [km]	M [$M_\odot$]	A [10^{57}]	E_b [10^{53}erg]
III-a	1.046	2.445	10.498	2.226	3.227	8.721	0.479	0.861	11.587	1.402	1.898	3.503
III-b	1.045	2.444	10.547	2.223	3.216	8.557	0.471	0.861	11.772	1.402	1.895	3.453
III-c	1.037	2.424	10.616	2.221	3.200	8.397	0.460	0.832	11.953	1.402	1.887	3.339
III-d	1.047	2.452	10.494	2.221	3.213	8.598	0.481	0.867	11.619	1.402	1.893	3.422
III-e	1.044	2.440	10.554	2.218	3.203	8.495	0.473	0.858	11.809	1.403	1.890	3.384
III-f	1.040	2.433	10.609	2.216	3.189	8.311	0.464	0.842	11.992	1.403	1.887	3.334
SLy230a [21]	1.15	2.69	10.25	2.10	2.99	7.07	0.508	0.925	11.8	1.4	1.85	2.60
SLy230b [21]	1.21	2.85	9.99	2.05	2.91	6.79	0.538	0.985	11.7	1.4	1.85	2.61

Non-spherical nucleons

The nucleon in a nucleus will include

- Local density approach for environment
- R dependence of a results
- Deformations
  - In particular, axially symmetric case allows the deformations in polar direction
- Polar deformations can be represented
  - in the isotopic vector and
  - in the profile function in ordinary space



$$\boldsymbol{N}(\boldsymbol{r}-\boldsymbol{R}) = \begin{pmatrix} \sin\Theta(\boldsymbol{r}-\boldsymbol{R})\cos\varphi \\ \sin\Theta(\boldsymbol{r}-\boldsymbol{R})\sin\varphi \\ \cos\Theta(\boldsymbol{r}-\boldsymbol{R}) \end{pmatrix}$$

$$P = P(|\boldsymbol{r}-\boldsymbol{R}|,\theta), \qquad \Theta = \Theta(|\boldsymbol{r}-\boldsymbol{R}|,\theta)$$

$$U(\boldsymbol{r}-\boldsymbol{R}) = \exp\left[i\boldsymbol{\tau}\cdot\boldsymbol{N}(\boldsymbol{r}-\boldsymbol{R})P(\boldsymbol{r}-\boldsymbol{R})\right]$$

Non-spherical nucleons

The Equations of Motion

- The coupled partial differential equations (not an easy problem)

$$f(F_{\tilde{r}\tilde{r}}, F_{\theta\theta}, F_{\tilde{r}}, F_{\theta}, \Theta_{\theta}, F, \Theta) = 0,$$
$$g(\Theta_{\theta\theta}, \Theta_{\theta}, F_{\tilde{r}}, F_{\theta}, \Theta, F) = 0,$$

- A numerical variational method can be applied

$$P(r,\theta) = 2\arctan\left\{\frac{r_0^2}{r^2}(1+m_\pi r)\big(1+u(\theta)\big)\right\} e^{-f(r)r}$$

$$\Theta(r,\theta) = \theta + \zeta(r,\theta),$$

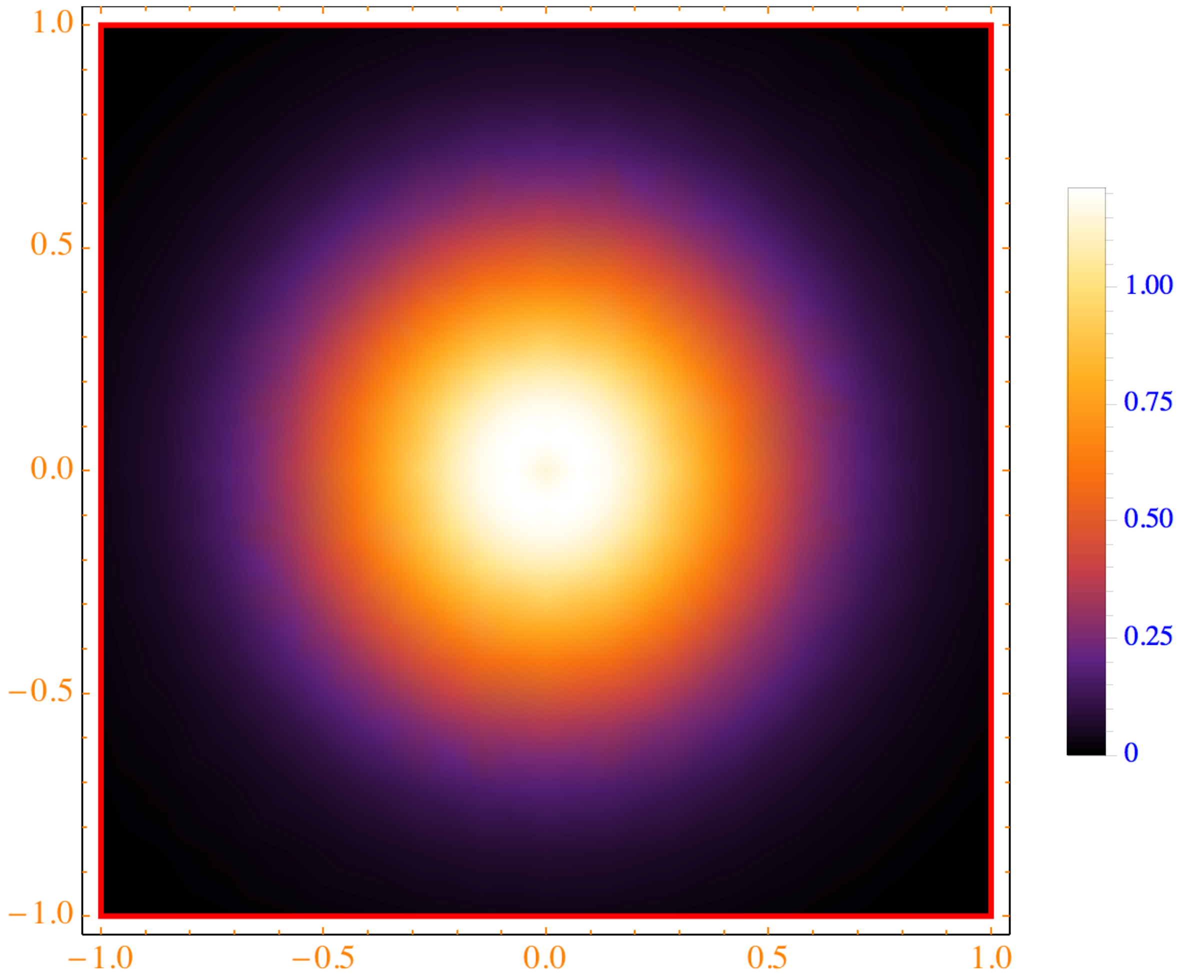
$$F(r) = 2\arctan\left\{\frac{r_0^2}{r^2}(1+m_\pi r)\right\} e^{-f(r)r}, \qquad u(\theta) = \sum_{n=1}^{\infty} \gamma_n \cos^n\theta$$

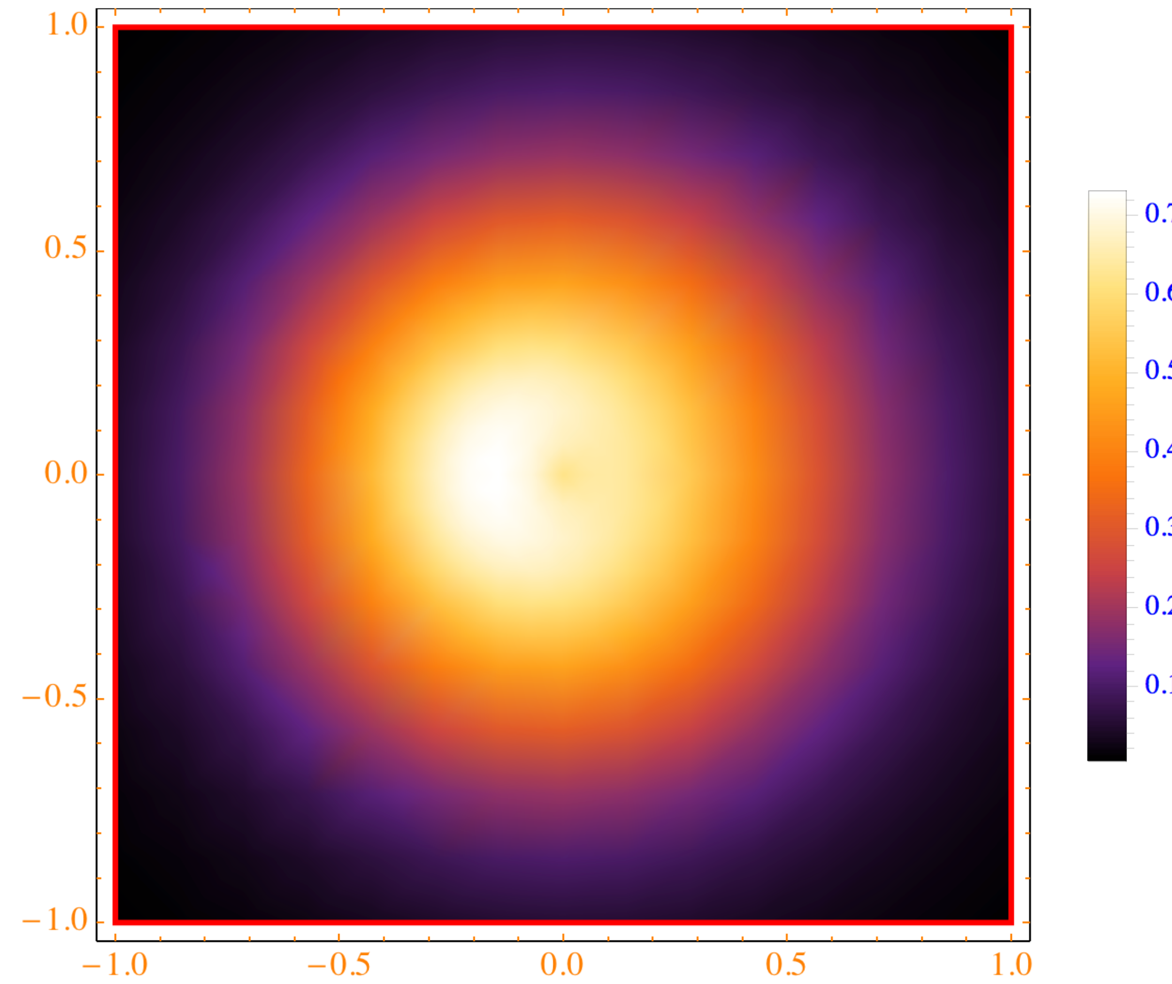
$$f(r) = \beta_0 + \beta_1 e^{\beta_2 r^2}.$$

$$\zeta(r,\theta) = r e^{-\delta_0^2 r^2} \sum_{n=1}^{\infty} \delta_n \sin 2n\theta,$$

$$\lim_{r\to 0} F(r) = \pi - Cr,$$
$$\lim_{r\to\infty} F(r) = D\,(1+m_\pi r)\,\frac{e^{-m_\pi r}}{r^2},$$

Baryon charge distribution inside the nucleon under the consideration

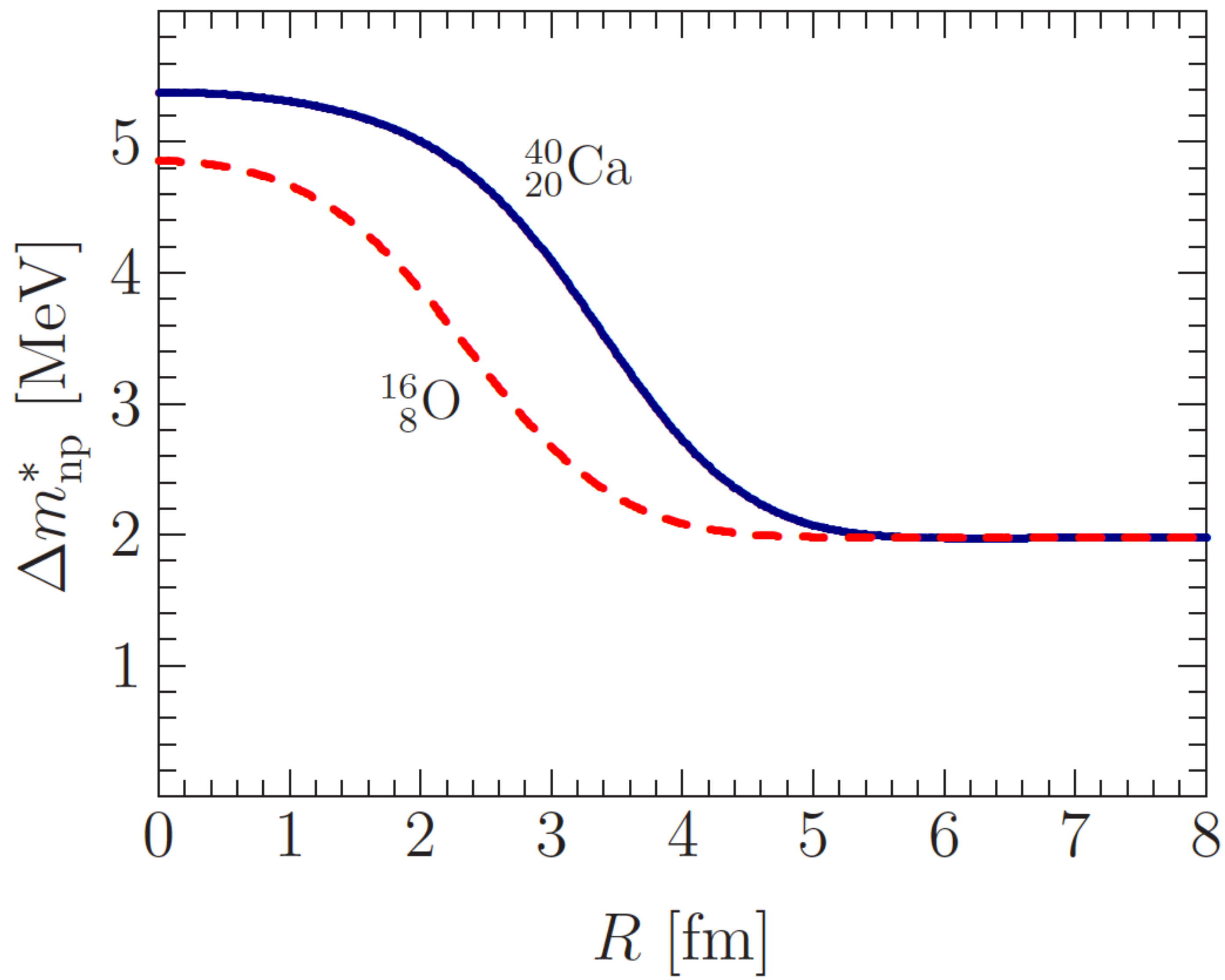




In free space (left) and in O16 (right), R = 1.5fm

Non-spherical nucleons

The neutron-proton mass difference in finite nuclei



R is a distance between the geometrical centres of nucleus and nucleon

Thank you very much for your attention!