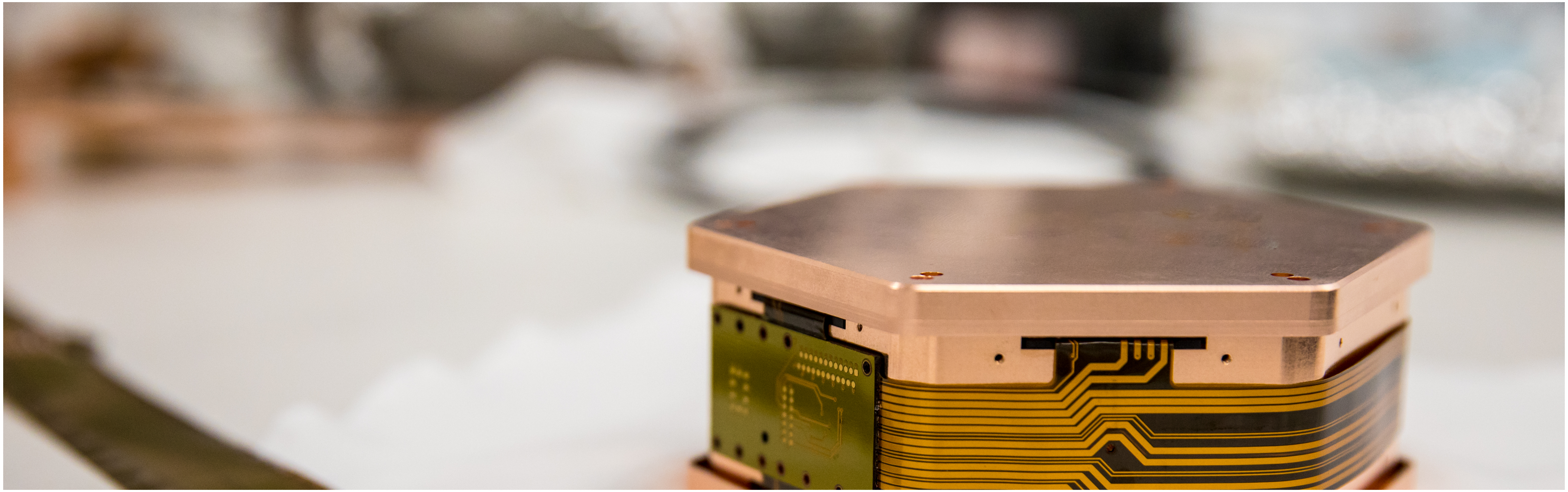


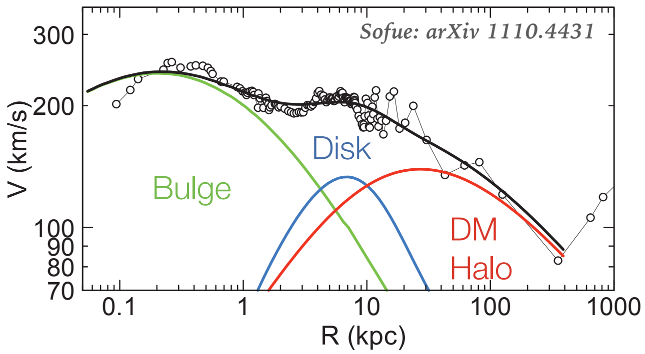
DARK MATTER DIRECT DETECTION OF CLASSICAL WIMPS

Jodi Cooley
SMU

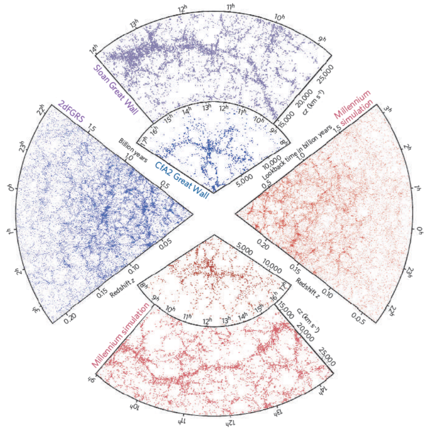
ABUNDANCE OF EVIDENCE FOR PARTICLE DARK MATTER

- The Missing Mass Problem:
  - Dynamics of stars, galaxies, and clusters
  - Rotation curves, gravitational lensing
  - Large Scale Structure formation
- Wealth of evidence for a particle solution
  - Microlensing (MACHOs) mostly ruled out
  - MOND has problems with Bullet Cluster
- Non-baryonic
  - Height of acoustic peaks in the CMB (Ω_b, Ω_m)
  - Power spectrum of density fluctuations (Ω_m)
  - Primordial Nucleosynthesis (Ω_b)
- And STILL HERE!
  - Stable, neutral, non-relativistic
  - Interacts via gravity and (maybe) a weak force

Rotation Curve of Milky Way

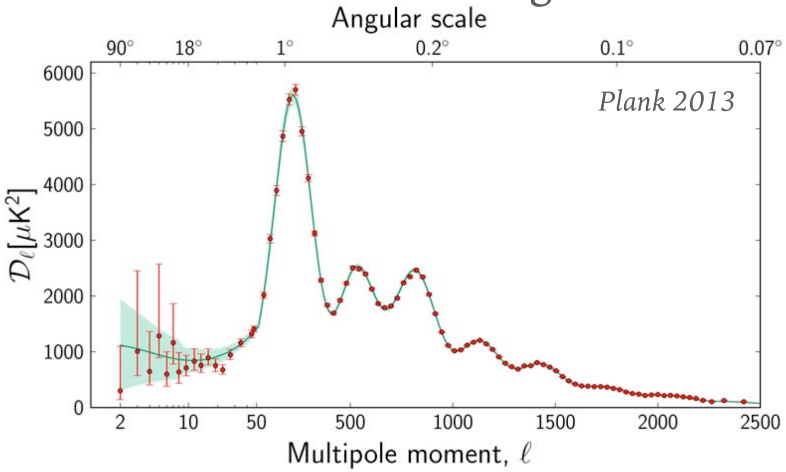


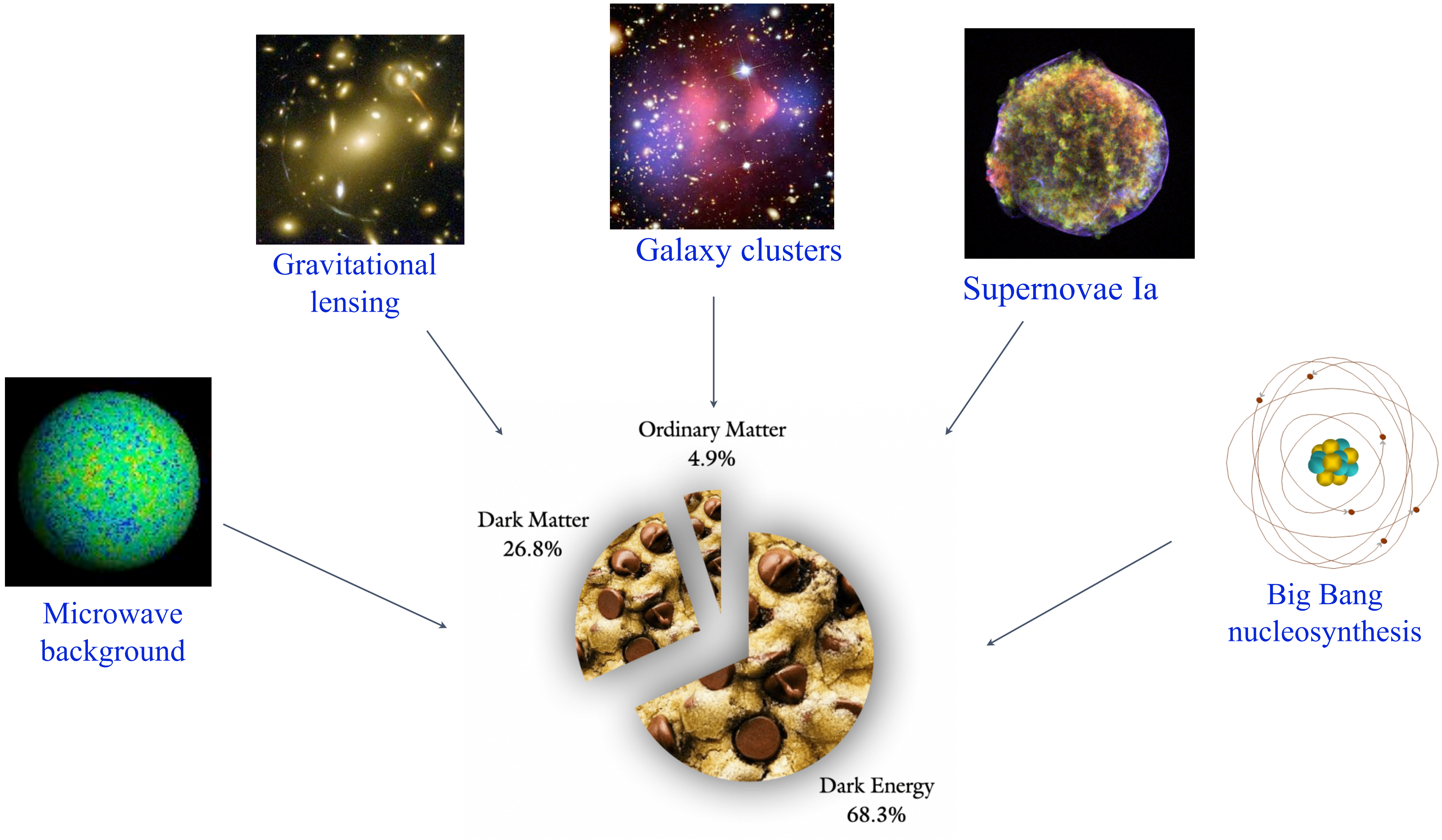
Structure Formation



Nature 2006

Cosmic Microwave Background

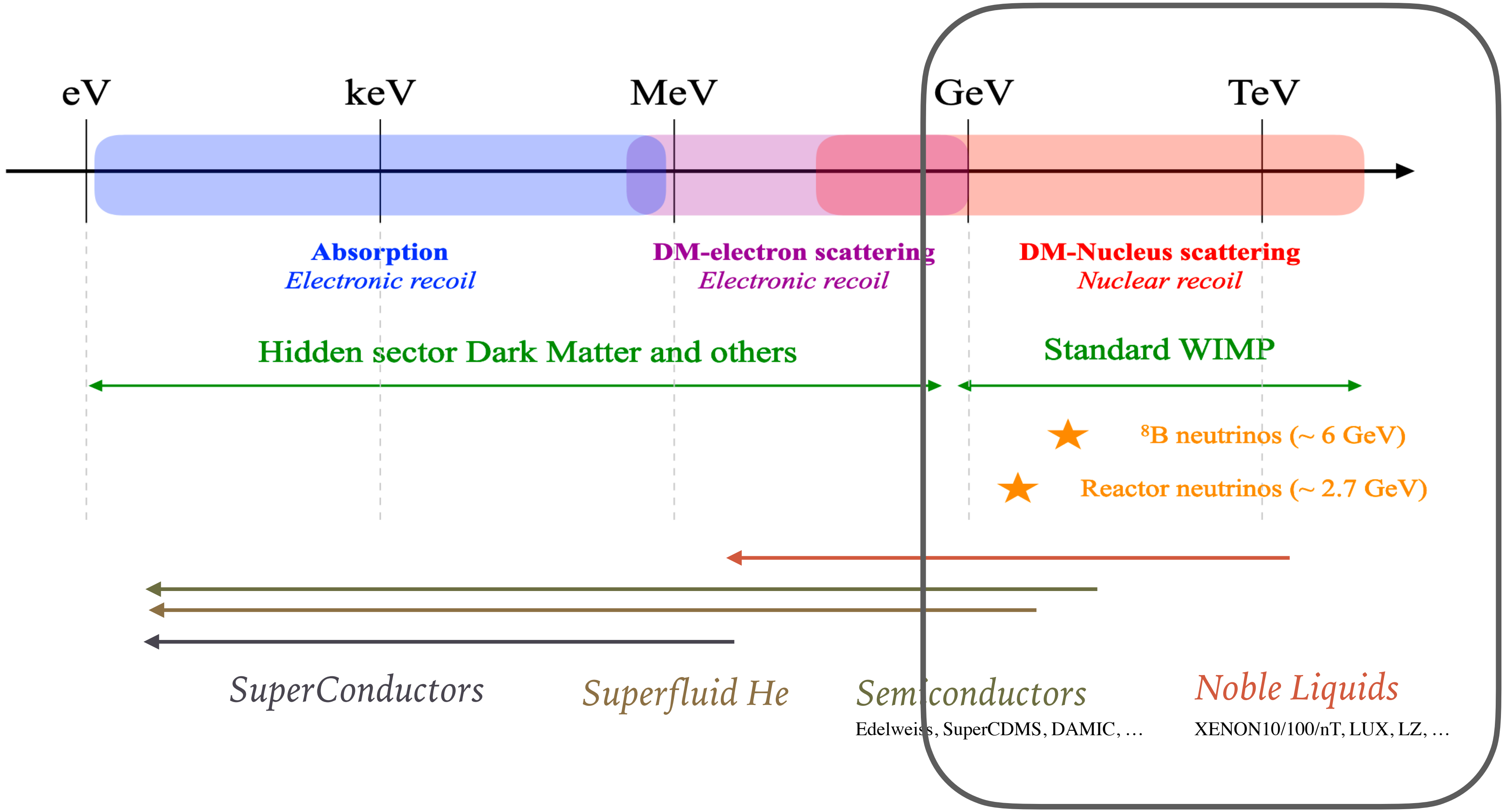



Gravitational lensing
Galaxy clusters
Supernovae Ia
Microwave background
Big Bang nucleosynthesis
Ordinary Matter 4.9%
Dark Matter 26.8%
Dark Energy 68.3%

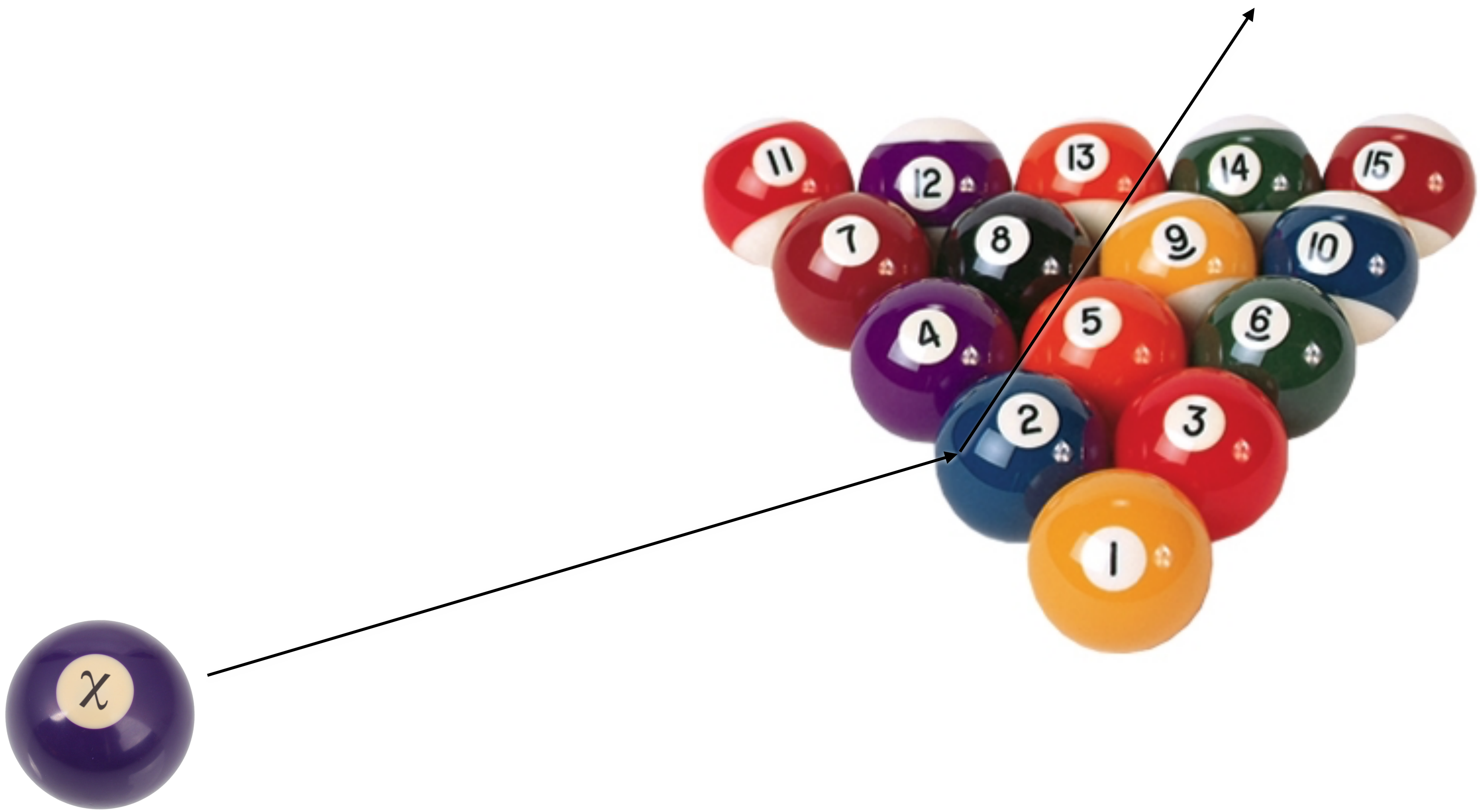
CONTENT

- How to Design a Dark Matter Detector
  - Expected rates
  - Background considerations
  - Experimental signatures
- Direct Detection Searches
  - Have we already seen a signal?
  - Detecting scattering from the nucleus with existing experiments
  - Reaching lower masses by detecting single electron-hole pairswith current experiments
  - Ideas for extending sensitivity to sub-eV dark matter signatures (will be completely covered by Tongyan later this week).

DIRECT DETECTION ENERGY RANGES

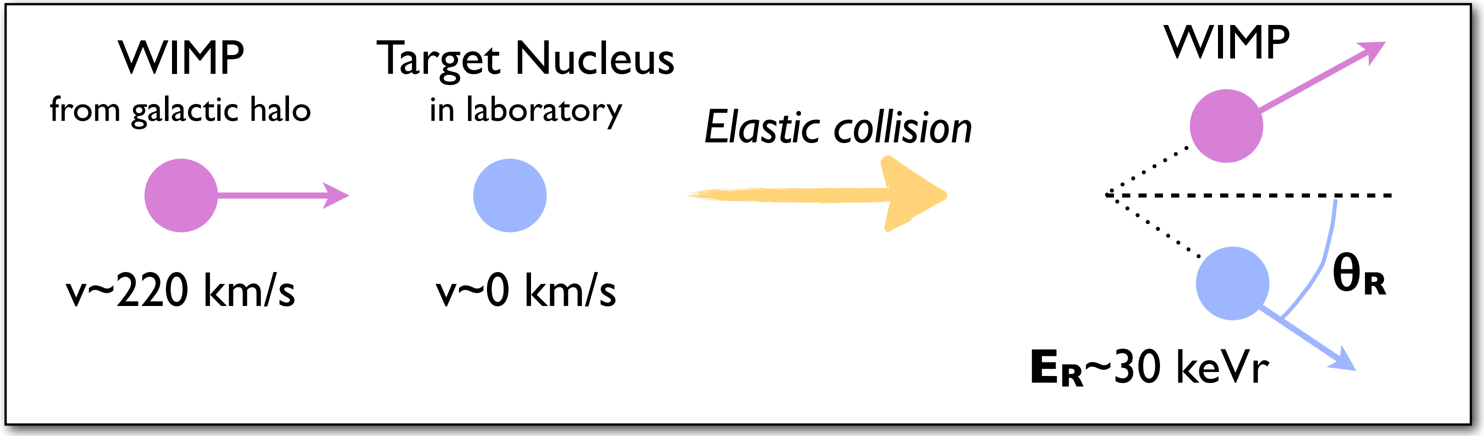


CONSIDERATIONS - DETECTING DARK MATTER VIA NUCLEAR SCATTERING



DIRECT DETECTION EVENT RATES

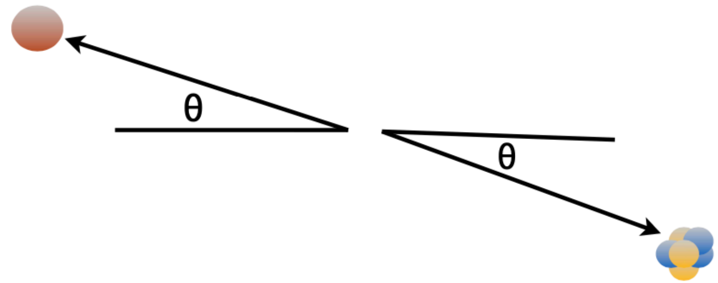
Assume that the dark matter is not only gravitationally interacting (WIMP).



- Elastic scatter of a WIMP off a nucleus
  - Imparts a small amount of energy in a recoiling nucleus
  - Can occur via spin-dependent or spin-independent channels
  - Need to distinguish this event from the overwhelming number of background events

DIVE-IN: KINEMATICS

- Calculate the recoil energy of a nucleus in the center of mass frame.



initial momentum: $\vec{p} = -\overrightarrow{E_k}$

final momentum: $\vec{p}\,' = -\overrightarrow{E_k}' = \vec{q} + \mu\vec{v}_\chi$

where

WIMP-nucleus reduced mass: $\mu = \dfrac{m_\chi m_N}{m_\chi + m_N}$

q = *momentum transfer*

- For elastic scattering in the COM frame: $|\vec{p}\,| = |\overrightarrow{p'}|$

$$\frac{q^2}{2} = \frac{1}{2}(\vec{p} - \vec{p}\,')^2 = p^2 - \vec{p}\cdot\vec{p}\,' = p^2(1-\cos\theta) = \mu^2 v^2(1-\cos\theta_R)$$

v = *mean WIMP-velocity relative to the target*

- The NR energy can then be calculated as

$$E_r = \frac{|\vec{q}\,|^2}{2m_N} = \frac{\mu^2 v^2}{m_N}(1-\cos\theta_R)$$

- We can calculate the minimum DM particle velocity for which we can expect a recoil. This corresponds to the case of backscattering ($\cos\theta = -1$).

$$E_R = \frac{\mu^2 v^2}{m_N}(1-\cos\theta_R)$$

$$E_r = \frac{|\vec{q}\,|^2}{2m_N}$$

$$E_R = \frac{\mu^2 v^2}{m_N}(1-(-1)) \Rightarrow \boxed{v_{min} = \sqrt{\frac{m_N E_R}{2\mu^2}} = \frac{q}{2\mu}}$$

- Implications:
  - Lighter dark matter particles ($m_\chi \ll m_N$) must have larger threshold velocities.
  - Inelastic scattering can further increase the minimal velocity needed.
- Consider the average momentum transfer in an elastic scattering between a WIMP-nucleus. Consider the case of a 10 GeV/c^2 WIMP whose speed is ~100 km s^{-1}.

$$p = m_\chi v = (10 \times 10^8\ eV\,c^{-2})(100 \times 10^3\ m\,s^{-1})\frac{c}{3 \times 10^8\ m\,s^{-1}} \sim 3\ MeV/c$$

If the DM were 100 GeV/c^2 then our momentum transfer would be ~ 30 MeV/c

- What is the de Broglie wavelength that corresponds to a momentum transfer of ~10 MeV/c?

$$\lambda = \frac{hc}{pc} = \frac{1.239 \times 10^{-6}\ eV \cdot m}{10 \times 10^{6}\ eV} \sim 12\ pm > R_0 A^{1/3}\ fm$$

This is larger than the size of most nuclei. Thus, scattering amplitudes on individual nucleons will add coherently.

EXPECTED RATES IN A DETECTOR - SIMPLIFIED.

- The differential event rate for simplified WIMP interaction (a detector stationary in the galaxy) is given by:

event rate · total event rate

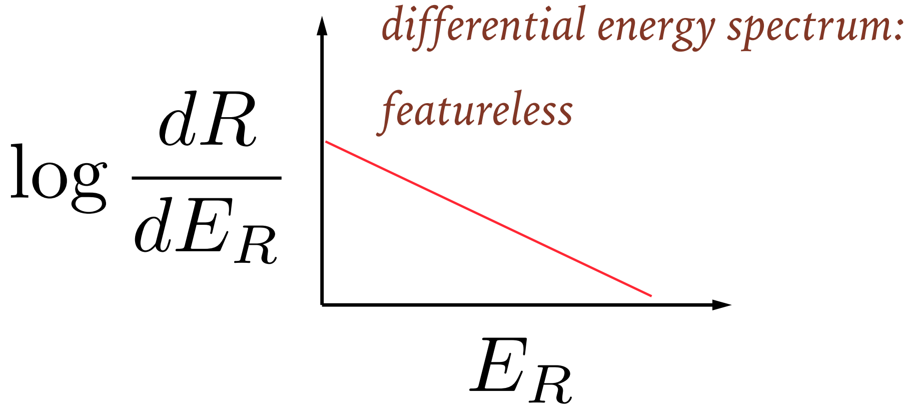
$$\frac{dR}{dE_R} = \frac{R_0}{E_0 r} e^{-E_R/E_0 r}$$

recoil energy

most probable incident energy (Maxwell-Boltzman distribution)

kinematic factor

$$\boxed{r = \frac{4M_\chi M_N}{(M_\chi + M_N)^2}}$$



- The total event rate is given by

$$\int_0^\infty \frac{dR}{dE_R} dE_R = R_0$$

and the mean recoiling energy

$$< E_R > = \int_0^\infty E_R \frac{dR}{dE_R} dE_R = E_0 r$$

EXAMPLE: CALCULATE THE MEAN NR DEPOSITED IN A DETECTOR

- Assume that the DM mass and the nucleus mass are identical:

$$m_\chi = m_N = 100\ GeV/c^2$$

- Our formula is

$$< E_R > = E_0 r \ = \left(\frac{1}{2} m_\chi v^2\right)\left(\frac{4 m_\chi m_N}{(m_\chi + m_N)^2}\right)$$

For our case:

$$\Longrightarrow r = 1$$

- Assuming the halo is stationary, the mean WIMP velocity relative to the target is

$$v \approx 220\ km\ s^{-1} = 0.75 \times 10^{-3}\ c$$

- Substituting into our equation for $<E_R>$

$$< E_R > = \frac{100\ GeV\ c^{-2}(0.75 \times 10^{-3}\ c)^2}{2} \ \sim 30\ keV$$

EXPECTED DETECTOR RATES: THE DETAILS

- We need to take into account the following
  - DM will have a certain velocity distribution $f(v)$.
  - The detector is on Earth, Earth moves around the Sun, and the Sun moves around the Galactic Center.
  - The cross-section depends upon the spin interaction. In the simplest cases, this is either spin-independent (SI) or spin-dependent (SD)
  - DM scatters on nuclei. Nuclei have finite size. As such, we have to consider form-factor corrections which are different for SI and SD interactions.
  - The recoil energy is not necessarily the observed energy. The detection efficiency in real life is not 100%.
  - Detectors have certain energy resolution and energy thresholds.

DARK MATTER DETECTION MASTER FORMULA

- The total number of particles detected (N) is the dark matter flux times the effective area of the target multiplied by the observation time (t)

$$N = t\underbrace{nv}_{\text{DM number density x DM speed}}\overbrace{N_T\sigma}^{\text{number of target x scattering cross section}}$$

- We will need to determine the spectrum of DM recoils → the energy dependence of the number of detected DM particles

$$\frac{dN}{dE_R} = tnvN_T\frac{d\sigma}{dE_R}$$

$$\frac{dN}{dE_R} = tnvN_T\frac{d\sigma}{dE_R}$$

➤ We need to consider the DM particles are described by their local velocity distribution, $f(\overrightarrow{v})$, were $\overrightarrow{v}$ is the DM velocity in the reference frame of the detector.

$$\frac{dN}{dE_R} = tnN_T\int_{v_{min}} vf(\overrightarrow{v})\frac{d\sigma}{dE_R}d\overrightarrow{v}$$

We need to integrate all possible DM velocities with their corresponding probability density and

$$v_{min} = \sqrt{\frac{m_\chi E_R}{2\mu^2}} =$$ *min speed required to produce a recoil of energy E_R.*

➤ Noting the following:

$$n = \frac{\rho}{m_\chi} \quad and \quad N_T = \frac{M_T}{m_N} \quad and \quad \epsilon = tM_T$$

where M_T is the total mass of the target and m_N is the mass of an individual nucleus

➤ We can write

$$\frac{dN}{dE_R} = \epsilon\frac{\rho}{m_\chi m_N}\int_{v_{min}} vf(\overrightarrow{v})\frac{d\sigma}{dE_R}d\overrightarrow{v}$$

ELEMENTS OF IDEAL EVENT RATE IN DIRECT DETECTION:

Differential Event Rate:
[events/keV/kg/day]

local WIMP density

WIMP-nucleon scattering cross section

$$\boxed{\frac{dR}{dE_R} = \frac{\rho_0}{m_N m_\chi}\int_{v_{min}}^{\infty} v f(v) \frac{d\sigma_{\chi N}}{dE_R} dv}$$

nucleus mass *WIMP mass* *WIMP speed distribution in detector frame*

need input from astrophysics, particle physics and nuclear physics

Minimum WIMP velocity which can cause a recoil of energy E_R.

$$v_{min} = \sqrt{\frac{m_N E_R}{2\mu^2}}$$

Elastic scattering happens in the extreme non-relativistic case in the lab frame.

$$\boxed{E_R = \frac{\mu_N^2 v^2 (1 - \cos\theta_R)}{m_N}}$$

$$\text{where } \mu = \frac{m_\chi m_N}{m_\chi + m_N} \text{ and } \theta_R = \text{scattering angle}$$

ELEMENTS OF IDEAL EVENT RATE IN DIRECT DETECTION:

Differential Event Rate:
[events/keV/kg/day]

local WIMP density

WIMP-nucleon scattering cross section

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nucleus mass

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$$\boxed{E_R = \frac{\mu_N^2 v^2 (1 - \cos\theta_R)}{m_N}}$$

where $\mu = \dfrac{m_\chi m_N}{m_\chi + m_N}$ *and* θ_R = *scattering angle*

THE SCATTERING CROSS SECTION

- Event rate is found by integrating over all recoils:

$$\boxed{R = \int_{E_T}^{\infty} dE_R \frac{\rho_0}{m_N m_\chi} \int_{v_{min}}^{\infty} v f(v) \frac{d\sigma_{\chi N}}{dE_R}(v, E_R) dv}$$

threshold energy

Minimum WIMP velocity which can cause a recoil of energy E_R.

$$v_{min} = \sqrt{\frac{m_N E_R}{2\mu^2}}$$

- The scattering cross section takes place in the non-relativistic limit. Thus, it can be approximated as isotropic.

$$\frac{d\sigma}{d\cos\theta} = constant = \frac{\sigma}{2}$$

$$\frac{d\sigma}{d\cos\theta} = \frac{\sigma}{2}$$

- Recall, $E_R^{max} = 2\mu^2 v^2 / m_N$. That means we can write …

$$E_R = E_R^{max}\frac{1+\cos\theta}{2} \quad \longrightarrow \frac{dE_R}{d\cos\theta} = \frac{E_R^{max}}{2}$$

- From this we can write

$$\frac{d\sigma}{dE_R} = \frac{d\sigma}{d\cos\theta}\frac{d\cos\theta}{dE_R} = \frac{\sigma}{2}\frac{2}{E_R^{max}} = \frac{\sigma}{E_R^{max}} = \frac{m_N}{2\mu}\frac{\sigma}{v^2}$$

- Recall that the momentum transfer for non-relativistic processes can neglect the kinetic energy of the nucleus. So,…

$$q = \sqrt{2m_N E_R} \sim MeV \qquad \rightarrow \textit{the de Broglie length is on the order of fm.}$$

- So, light nuclei, the DM particle sees the nucleus as a whole, w/o substructure.

- Heavier nuclei require inclusion of the nuclear form factor to account for the loss of coherence.

- The WIMP-nucleon cross section can be separated:

$$\frac{d\sigma}{dE_R} = \left[\left(\frac{d\sigma}{dE_R}\right)_{SI} + \left(\frac{d\sigma}{dE_R}\right)_{SD}\right]$$

Spin-Independent + Spin-Dependent

SI arise from scalar or vector couplings to quarks.

SD arise from axial-vector couplings to quarks.

- To calculate, add coherently the spin and scalar components

$$\boxed{\frac{d\sigma}{dE_R} = \frac{m_N}{2\mu_N^2 v^2}\left[\sigma_0^{SI} F_{SI}^2 + \sigma_0^{SD} F_{SD}^2\right]}$$

Particle Theory

Nuclear Structure

$F(E_R)$ = Form factor encodes the dependence on the momentum transfer.

$$\frac{d\sigma_{\chi N}}{dE_R} = \frac{m_N}{2\mu_N^2 v^2}\left[\sigma_0^{SI} F_{SI}^2 + \sigma_0^{SD} F_{SD}^2\right]$$

NUCLEAR FORM FACTOR - QUANTUM MECHANICS OF INTERACTION WITH NUCLEUS

- Spin Independent: Woods-Saxon Form Factor

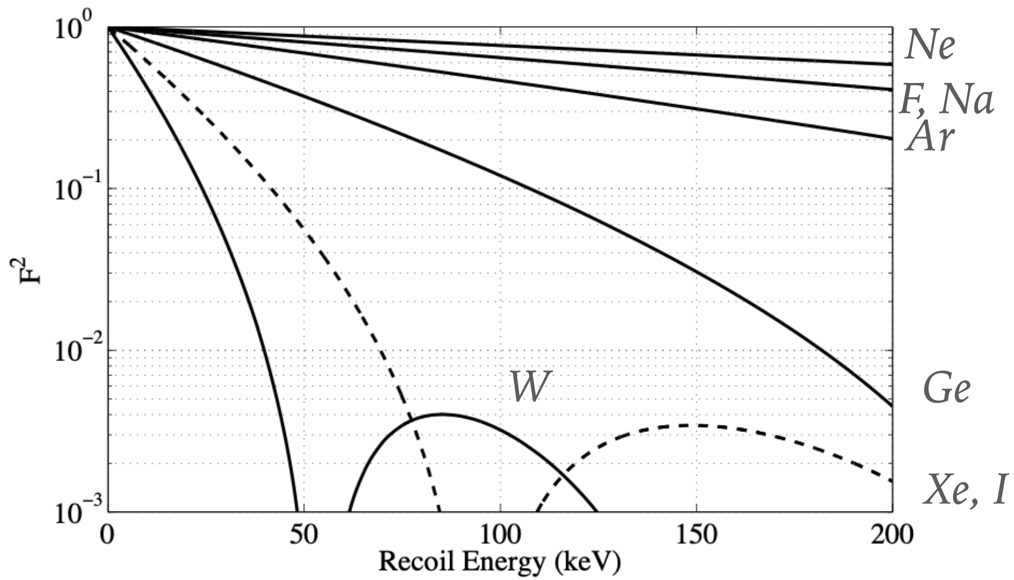
$$F(q) = \left(\frac{3j_1(qR_1)}{qR_1}\right)^2 e^{-q^2s^2/2}$$

$$j_1 = \textit{spherical Bessel Function} = \frac{\sin(x)}{x^2} - \frac{\cos(x)}{x}$$

q = *momentum transfer*

s = *nuclear skin thickness* ($\simeq$ *1 fm*)

R_1 = *effective nucleus radius*



$$\frac{d\sigma_{WN}}{dE_R} = \frac{m_N}{2\mu_N^2 v^2}\left[\sigma_0^{SI} F_{SI}^2 + \sigma_0^{SD} F_{SD}^2\right]$$

NUCLEAR FORM FACTOR - QUANTUM MECHANICS OF INTERACTION WITH NUCLEUS

- Spin Dependent Interactions

$$F^2(E_R) = \frac{S(E_R)}{S(0)}$$

$$S(E_R) = a_0^2 S_{00}(E_R) + a_1^2 S_{11}(E_R) + a_0 a_1 2S_{01}(E_R)$$

$$a_0 = a_p + a_n \text{ and } a_1 = a_p - a_n$$

$S_{ij} \longrightarrow$ *isoscalar, isovector and interference form factors*

$a_{i,j} \longrightarrow$ *isoscalar, isovector coupling*

$$\frac{d\sigma_{WN}}{dE_R} = \frac{m_N}{2\mu_N^2 v^2}\left[\sigma_0^{SI} F_{SI}^2 + \sigma_0^{SD} F_{SD}^2\right]$$

WIMP-NUCLEUS INTERACTIONS

- **Spin-Independent**

$$\sigma_0^{SI} = \frac{4\mu^2}{\pi}\Big[Z\underset{\substack{\text{coupling to}\\ \text{proton}}}{f_p} + (A-Z)\underset{\substack{\text{coupling to}\\ \text{neutron}}}{f_n}\Big]^2 \propto A^2$$

Assume low momentum transfer:

- In most models $f_n \sim f_p$ (scalar four-fermion coupling constants)
  - Scattering adds coherently with A^2 enhancement

- **Spin-Dependent**

Fermi constant

expectation value of proton/neutron spin

$$\sigma_0^{SD} = \frac{32 G_F^2 \mu^2}{\pi}\frac{J+1}{J}\Big[a_p < S_p > + a_n < S_n >\Big]^2$$

nuclear angular momentum

coupling constant to proton/neutron

- Scales with spin of the nucleus
  - No coherent effect!

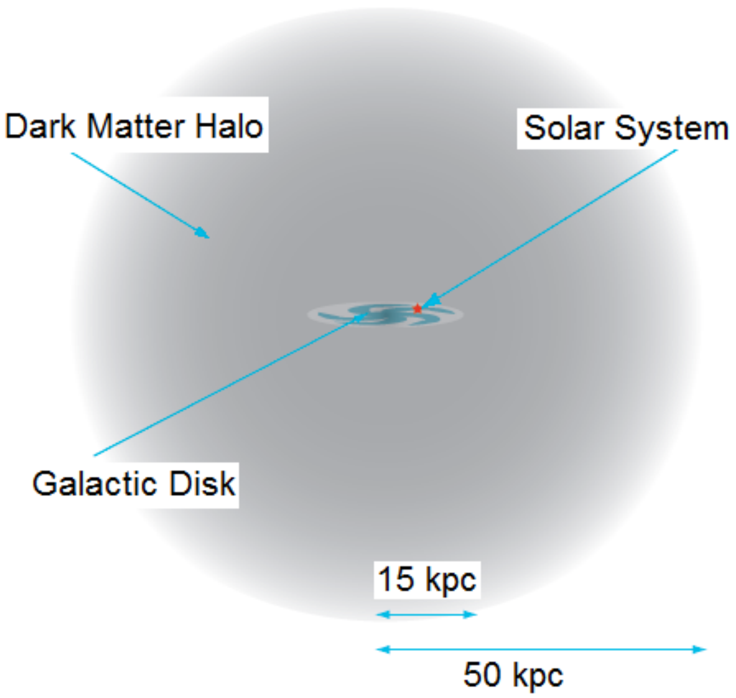
CHOOSING A TARGET MATERIAL: SPIN DEPENDENT CASE

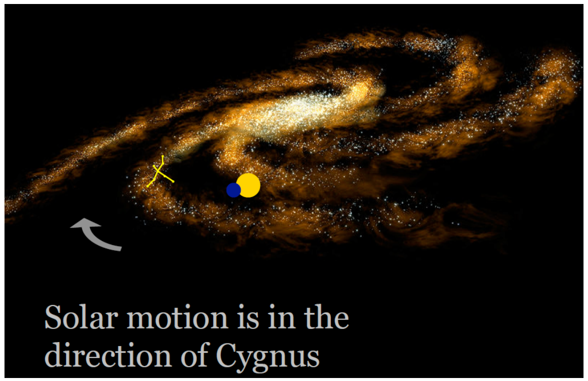
$$\sigma_0^{SD} = \frac{32G_F^2\mu^2}{\pi}\frac{J+1}{J}\left[a_p < S_p > + a_n < S_n >\right]^2$$

"Scaling Factors"

Nucleus	Z	Odd Nucleon	J	$\langle S_p \rangle$	$\langle S_n \rangle$	$\frac{4\langle S_p\rangle^2(J+1)}{3J}$	$\frac{4\langle S_n\rangle^2(J+1)}{3J}$
^{19}F	9	p	1/2	0.477	-0.004	9.10×10^{-1}	6.40×10^{-5}
^{23}Na	11	p	3/2	0.248	0.020	1.37×10^{-1}	8.89×10^{-4}
^{27}Al	13	p	5/2	-0.343	0.030	2.20×10^{-1}	1.68×10^{-3}
^{29}Si	14	n	1/2	-0.002	0.130	1.60×10^{-5}	6.76×10^{-2}
^{35}Cl	17	p	3/2	-0.083	0.004	1.53×10^{-2}	3.56×10^{-5}
^{39}K	19	p	3/2	-0.180	0.050	7.20×10^{-2}	5.56×10^{-3}
^{73}Ge	32	n	9/2	0.030	0.378	1.47×10^{-3}	2.33×10^{-1}
^{93}Nb	41	p	9/2	0.460	0.080	3.45×10^{-1}	1.04×10^{-2}
^{125}Te	52	n	1/2	0.001	0.287	4.00×10^{-6}	3.29×10^{-1}
^{127}I	53	p	5/2	0.309	0.075	1.78×10^{-1}	1.05×10^{-2}
^{129}Xe	54	n	1/2	0.028	0.359	3.14×10^{-3}	5.16×10^{-1}
^{131}Xe	54	n	3/2	-0.009	-0.227	1.80×10^{-4}	1.15×10^{-1}

Tovey et al., PLB 488 17(2000)





RELIC WIMP DISTRIBUTION: SIMPLIFIED MODEL

- WIMPs are distributed in isothermal spherical halos with Gaussian velocity distribution (Maxwellian)

$$f(\vec{v}) = \frac{1}{\sqrt{2\pi\sigma}} e^{-\frac{|\vec{v}|^2}{2\sigma^2}}$$

- The speed dispersion is related to the local circular speed by

$$\sigma = \sqrt{\frac{3}{2}} v_c \quad where \quad v_c = 220 \ km/s$$

- The density profile of the sphere is

$$\rho(r) \propto r^{-2} \quad and \quad \rho_0 = 0.3 \ GeV/c^2$$

- Particles with speeds greater than v_{esc} are not gravitationally bound. Hence, the speed distribution needs to be truncated.

$$v_{esc} = 650 \ km/s$$

DENSITY OF WIMPS IN YOUR WORK AREA

- The local dark matter density is

$$\rho_0 = 0.3\ GeV/cm^3$$

- Pick your favored mass for the dark matter particle

$$m = 5\ GeV/c^2$$

$$m = 60\ GeV/c^2$$

- What is the number density?

$$60{,}000\ particles/m^3 \longrightarrow for\ 5\ GeV/c^2$$

$$5{,}000\ particles/m^3 \longrightarrow for\ 60\ GeV/c^2$$

- How many dark matter particles in a 2 liter bottle?

$$recall\ that\ 1\ liter = 0.001\ m^3$$

$$120\ particles \longrightarrow for\ 5\ GeV/c^2$$

$$10\ particles \longrightarrow for\ 60\ GeV/c^2$$

Joachim Brod
Week 4!

MAYBE NOT THAT SIMPLE?

- Effective Field Theory considers leading order and NLO operators that can occur in the effective Lagrangian that describes the WIMP-nucleon interactions.
- Contains 14 operators, that rely on a range of nuclear properties in addition to the SI and SD cases. They combine such that the WIMP-nucleon cross section depends on six independent nuclear response functions:
  - One "Spin independent"
  - Two "Spin Dependent"
  - Three "Velocity-Dependent"
- Two pairs of these interfere, resulting in eight independent parameters that can be probed

The effective field theory of dark matter direct detection

A. Liam Fitzpatrick,[a] Wick Haxton,[b] Emanuel Katz,[a,c,d] Nicholas Lubbers,[c] Yiming Xu[c]

http://arxiv.org/abs/1211.2818

http://arxiv.org/abs/1308.6288

http://arxiv.org/abs/1405.6690

http://arxiv.org/abs/1503.03379

$$\mathcal{O}_1 = 1_\chi 1_N$$
$$\mathcal{O}_3 = i\vec{S}_N \cdot \left[\frac{\vec{q}}{m_N} \times \vec{v}^\perp\right]$$
$$\mathcal{O}_4 = \vec{S}_\chi \cdot \vec{S}_N$$
$$\mathcal{O}_5 = i\vec{S}_\chi \cdot \left[\frac{\vec{q}}{m_N} \times \vec{v}^\perp\right]$$
$$\mathcal{O}_6 = \left[\vec{S}_\chi \cdot \frac{\vec{q}}{m_N}\right]\left[\vec{S}_N \cdot \frac{\vec{q}}{m_N}\right]$$
$$\mathcal{O}_7 = \vec{S}_N \cdot \vec{v}^\perp$$
$$\mathcal{O}_8 = \vec{S}_\chi \cdot \vec{v}^\perp$$
$$\mathcal{O}_9 = i\vec{S}_\chi \cdot \left[\vec{S}_N \times \frac{\vec{q}}{m_N}\right]$$
$$\mathcal{O}_{10} = i\vec{S}_N \cdot \frac{\vec{q}}{m_N}$$
$$\mathcal{O}_{11} = i\vec{S}_\chi \cdot \frac{\vec{q}}{m_N}$$
$$\mathcal{O}_{12} = \vec{S}_\chi \cdot \left[\vec{S}_N \times \vec{v}^\perp\right]$$
$$\mathcal{O}_{13} = i\left[\vec{S}_\chi \cdot \vec{v}^\perp\right]\left[\vec{S}_N \cdot \frac{\vec{q}}{m_N}\right]$$
$$\mathcal{O}_{14} = i\left[\vec{S}_\chi \cdot \frac{\vec{q}}{m_N}\right]\left[\vec{S}_N \cdot \vec{v}^\perp\right]$$
$$\mathcal{O}_{15} = -\left[\vec{S}_\chi \cdot \frac{\vec{q}}{m_N}\right]\left[\left(\vec{S}_N \times \vec{v}^\perp\right) \cdot \frac{\vec{q}}{m_N}\right]$$

- The EFT framework parameterizes the WIMP-nucleus interaction in terms of the 14 operators listed to the left.

  $\overrightarrow{v}^\perp$ = relative velocity between incoming WIMP and nucleon

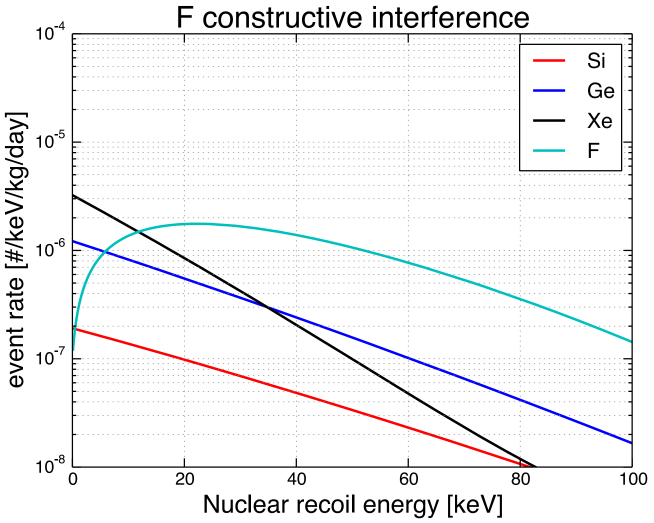
  q = momentum transfer

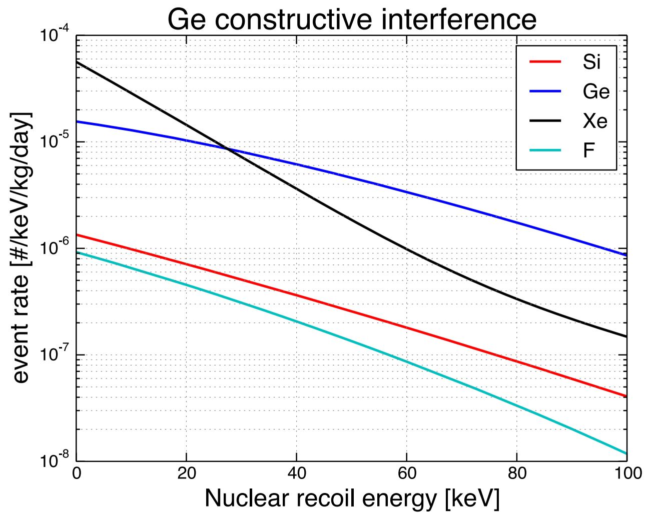
  $\overrightarrow{S}_\chi$ = WIMP spin

  $\overrightarrow{S}_N$ = nucleon spin

- In addition, each operator can independently couple to protons or neutrons.

  Note $\mathcal{O}_2$ is not considered as it cannot arise from the non-relativistic limit.





arxiv: 1503.03379

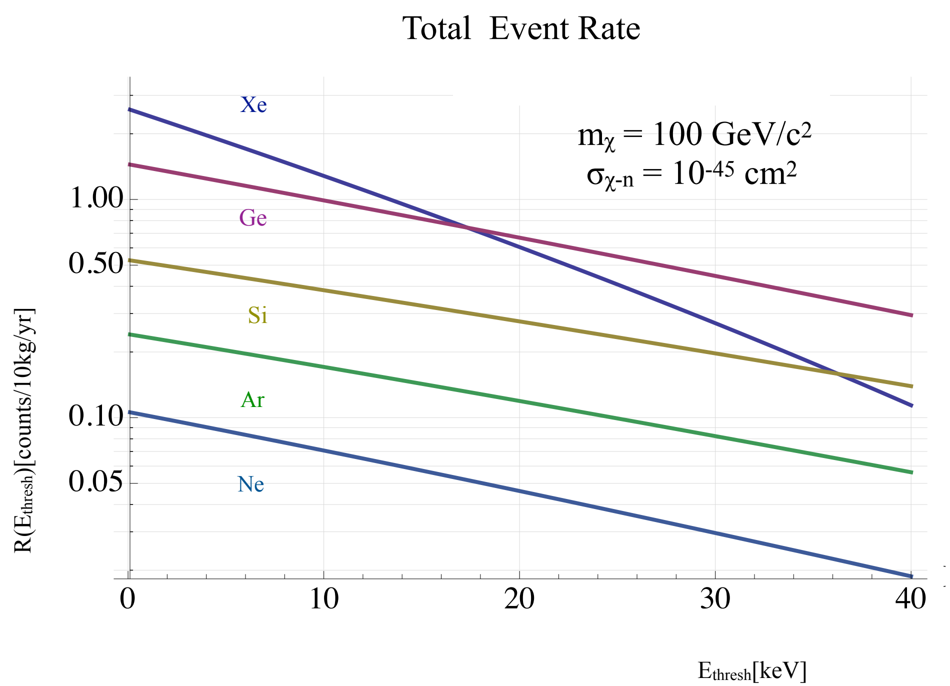
DARK MATTER COULD LOOK DIFFERENT IN DIFFERENT TARGETS

- Nuclear responses for different target elements vary. Some EFT operations have momentum dependance. EFT Operators can interfere.
  - Example illustrates differences evaluating at the $\mathcal{O}_8$ and $\mathcal{O}_9$ constructive interference vector.
  - Results in different rates between targets AND different spectral shapes.
- A robust dark matter direct detection program with different target materials will be needed to nail down which operators are contributing to any detected signal
- Take home message: We will need multiple targets to map out the physics of WIMP-nucleon interactions!

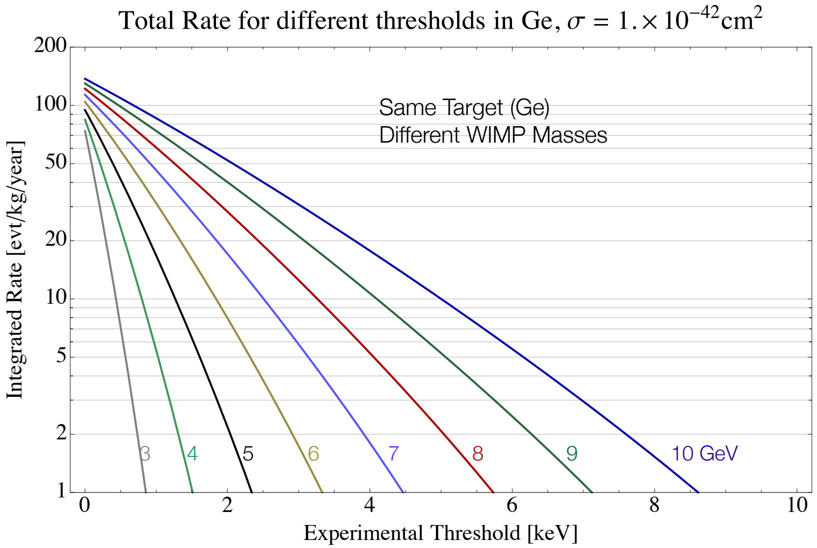
More on this from Joachim Brod during Week 4!

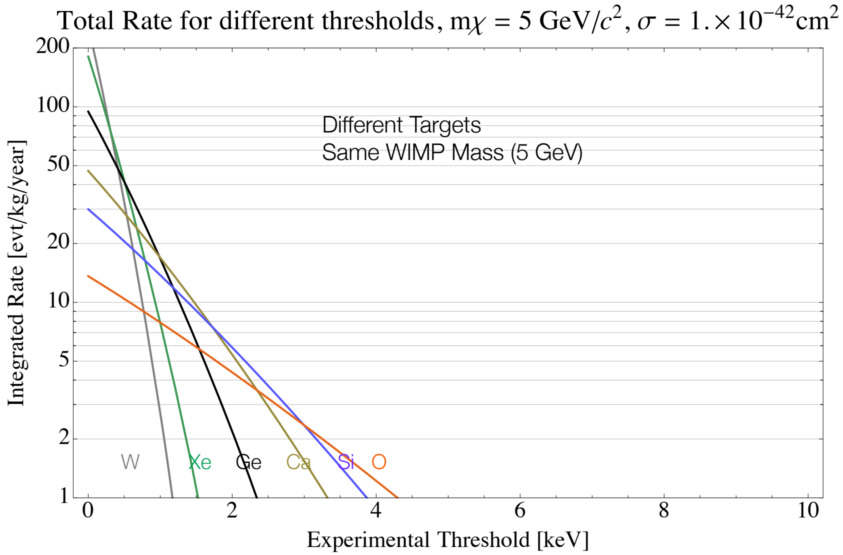
EVENT RATES ARE EXTREMELY LOW!

- Elastic scattering of WIMP deposits small amounts of energy into a recoiling nucleus (~few 10s of keV)
- Featureless exponential spectrum with no obvious peak, knee, break …
- Event rate is very, very low.
- Radioactive background of most materials is higher than the event rate.

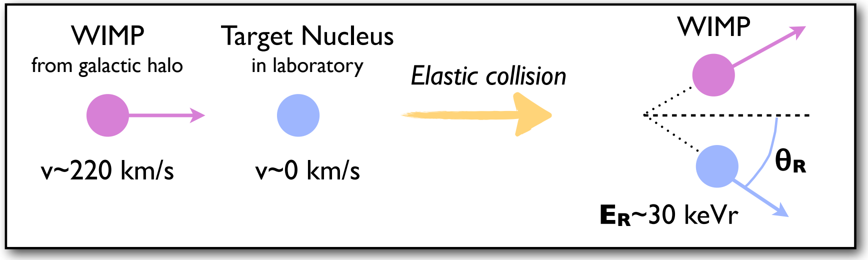


Need large exposures (mass x time)!



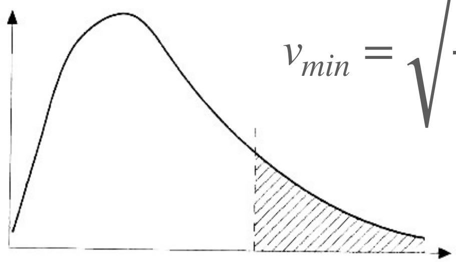


THE LOW-MASS WIMP CHALLENGE



$$E_R = \frac{p^2}{2m_N} = \frac{m_\mu^2 v^2}{m_N}(1 - \cos\theta_R)$$

A WIMP must have a minimum velocity to produce a recoil.

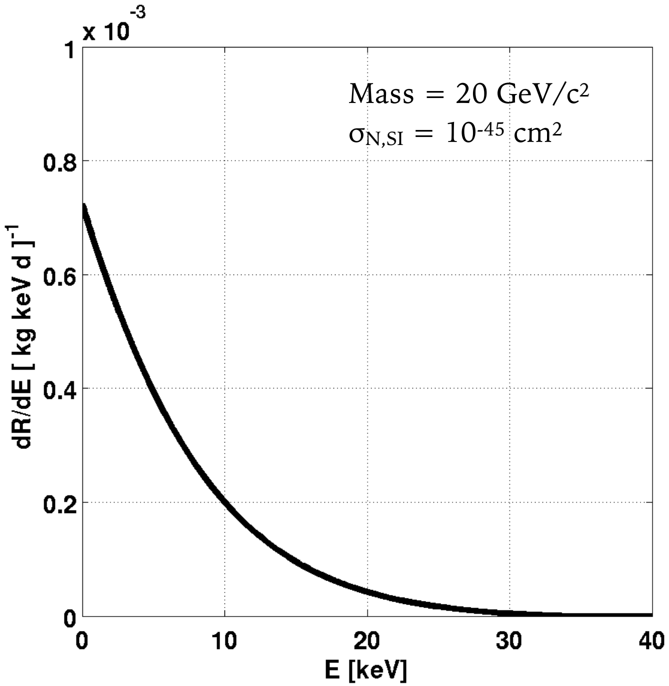


$$v_{min} = \sqrt{\frac{m_N E_{th}}{2m_\mu^2}}$$

Need Low Energy Threshold!

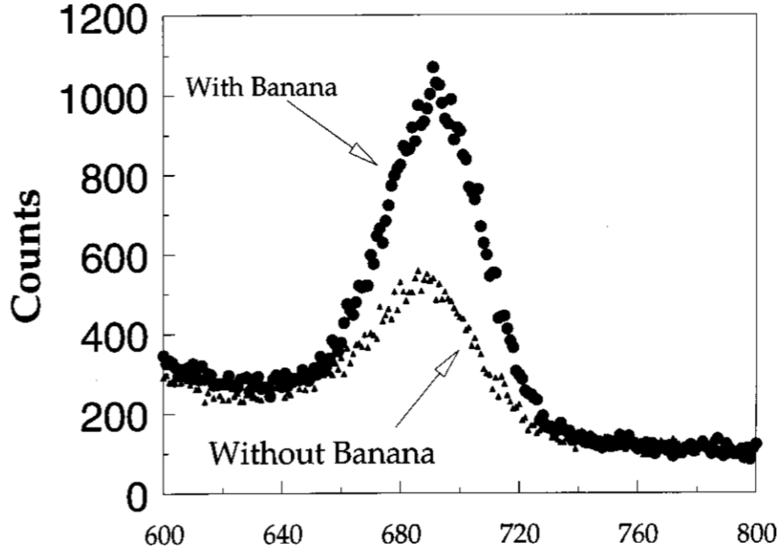
THE EVENT RATES ARE EXTREMELY LOW!

- Expected WIMP Spectrum



- Measured Banana Spectrum

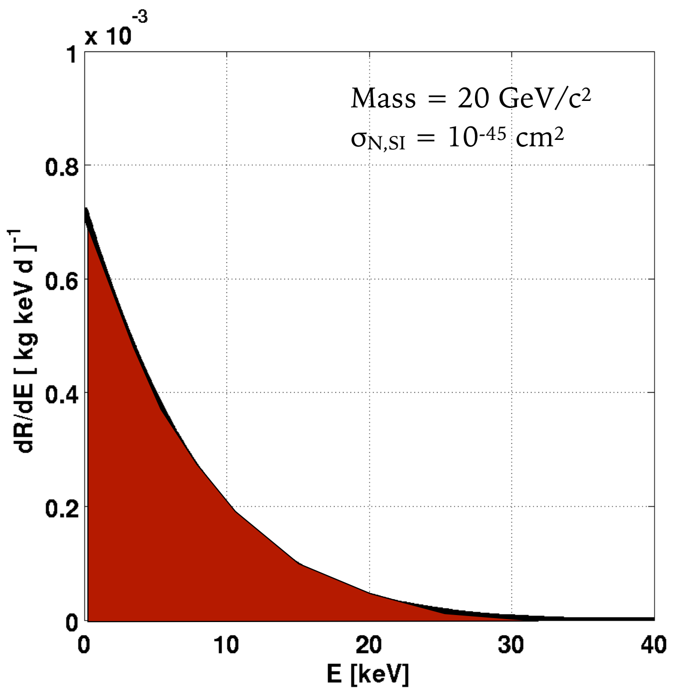
Hoeling et al Am.J.Phys. 1999, 67, 440.



Gamma measurements with a 3-inch NaI detector

THE EVENT RATES ARE EXTREMELY LOW!

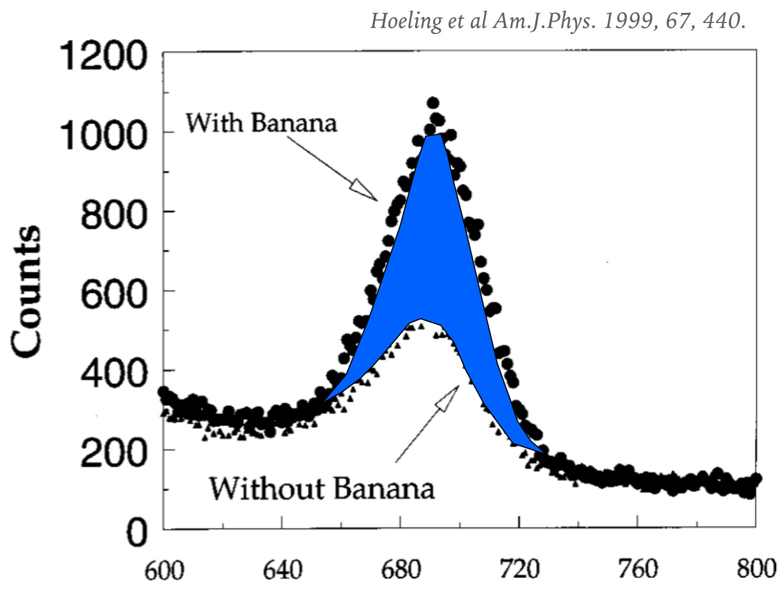
- Expected WIMP Spectrum



~1 event per kg per year

(nuclear recoils)

- Measured Banana Spectrum



~100 events per kg per year

(electron recoils)