CMB Probe of Particle Dark Matter and Primordial Black Hole

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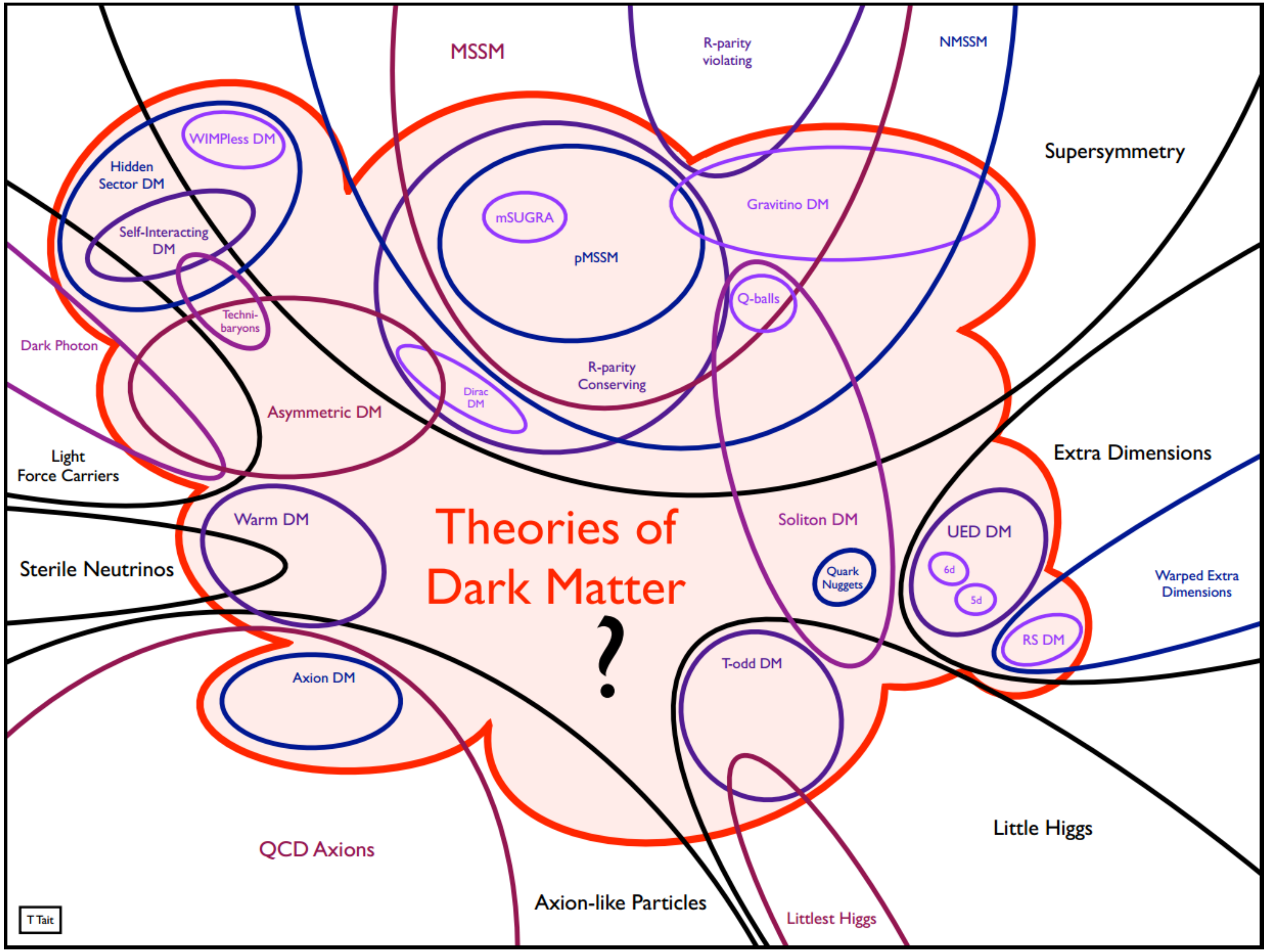
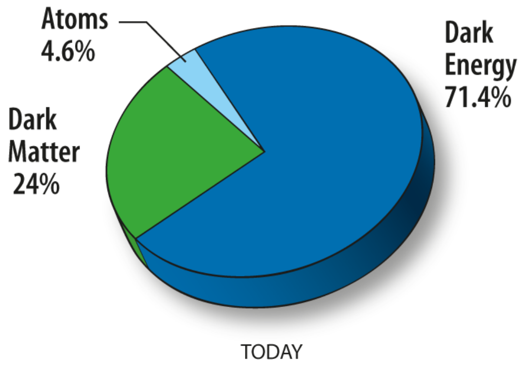
ArXiv 2002.03380 2011.12244

PowerPoint

What is Dark Matter





What we know

1. Gravitational Interactive
2. Negligible interaction with Baryonic Matter
3. Stable
4. Dynamically Cold

Possibilities

1. New Particle(s)
2. Primordial Black Hole (PBH)
3. Modified Gravity

. . .

PBH

Important probe of the very early universe !

Formation

(Very early universe)

1. Density Fluctuation
2. Inflationary Perturbation
3. Cosmological Phase Transition
4. Collapsing Domain Wall Bubble
5. Collapsing Cosmic String

. . .

Mass Range

$$M \sim \frac{c^3\,t}{G} \sim 10^{15} \left(\frac{t}{10^{-23}\,\mathrm{s}} \right) \mathrm{g}.$$

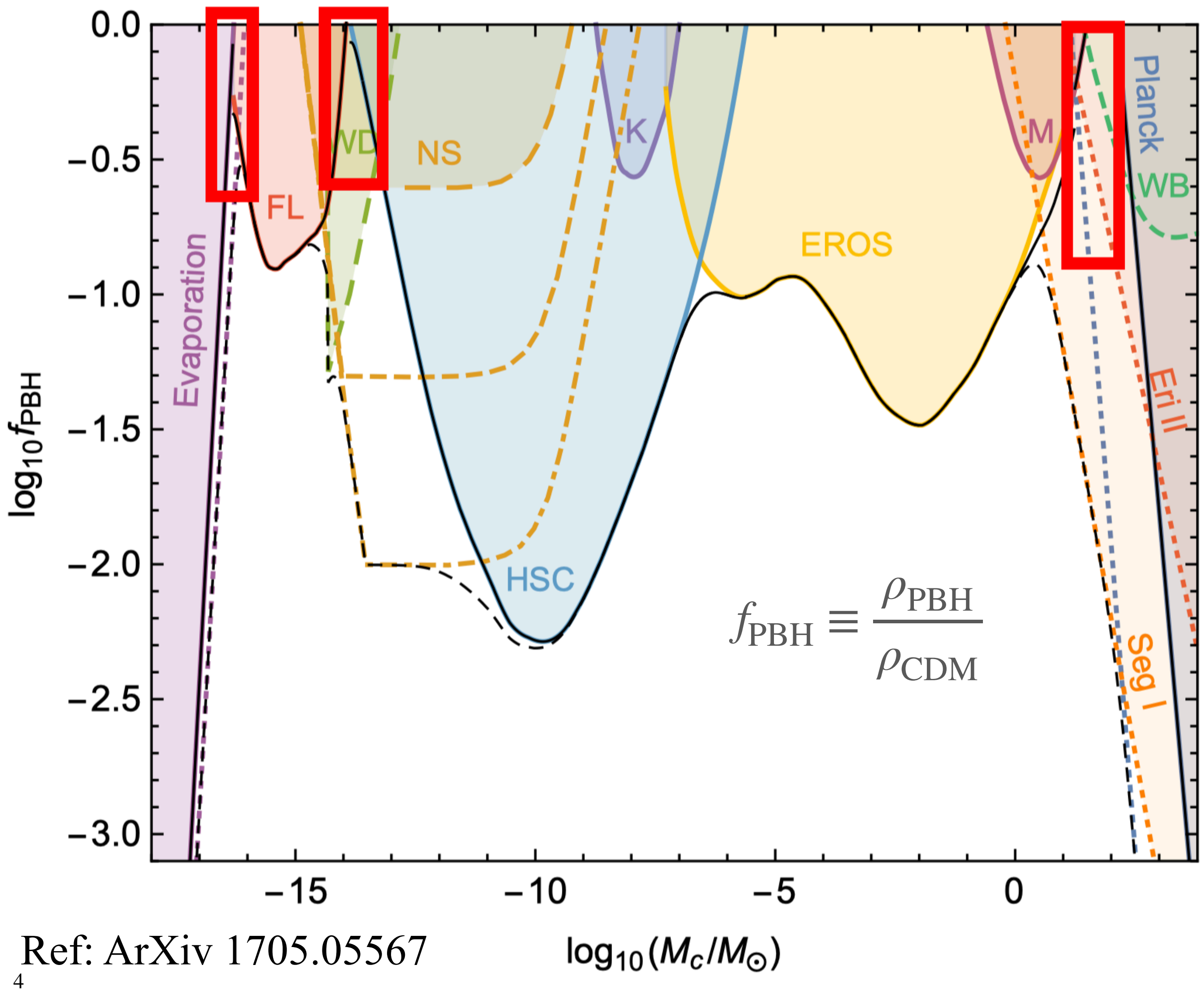
Anywhere between

$$10^{-5}\mathrm{g} \;\sim\; 10^5 M_\odot$$

Theories

- Quantum Gravity
- Inflation

PBH Current Status



Mechanism

- Gravitational Lensing
- Gravitational Wave
- Accretion
- Hawking Radiation
- …

Only 3 Mass Windows Left

Our Window: $10^{13} \sim 10^{17}$g

Ref: ArXiv 1705.05567

Overview

Particle/Energy Injection

Dark Matter (DM)

- Annhilation: $\chi\bar{\chi} \rightarrow S\bar{S}$
- Decay: $\chi \rightarrow S\bar{S}$

Primordial Black Hole (PBH):

Hawking Radiation

Energy Deposition

Ionization

- H Ionization
- H Excitation
- He Ionization

Heating

CMB Energy Spectra

Observables

x_e → CMB Anisotropy

21cm Signal

Can DM Ionize the Universe?

Energy required (unit comoving volume):

$$\Delta E = n_H \times 13.6eV = 13.6eV \cdot \frac{\Omega_b \rho_{cr}}{m_p}$$

DM density

$$\Delta \rho_c = \rho_c (1 - e^{-t/\tau}) = \rho_{cr} \Omega_c t / \tau$$

Assuming

$$\Delta E = f \times \Delta \rho_c$$

$$f \sim 0.1$$

Lower Bound:

$$\tau \sim 1.5 \times 10^{25} s$$

Upper bound:

$$\frac{\Delta \rho_c}{\rho_c} \sim \frac{t}{\tau} \sim 3 \times 10^{-8}$$

Dark Matter Injection

Dark Matter: $$\left(\frac{\mathrm{d}E}{\mathrm{d}V\mathrm{d}t}\right)_{\mathrm{INJ}} = R \times E_{\mathrm{n}}$$

Decay:

$$R^{\mathrm{dec}} = \Gamma_\chi n_\chi e^{-\Gamma_\chi t},\ E_{\mathrm{n}}^{\mathrm{dec}} = m_\chi$$

$$\left(\frac{\mathrm{d}E}{\mathrm{d}V\mathrm{d}t}\right)_{\mathrm{INJ}}^{\mathrm{dec}} = \Gamma_\chi \Omega_\chi (1+z)^3 \rho_{\mathrm{c}}$$

Annihilation (s-wave):

$$R^{\mathrm{ann}} = \mathrm{g}\ n_\chi^2 \langle \sigma v \rangle,\ E_{\mathrm{n}}^{\mathrm{ann}} = 2m_\chi$$

Homogeneous $$\left(\frac{\mathrm{d}E}{\mathrm{d}V\mathrm{d}t}\right)_{\mathrm{INJ}}^{\mathrm{ann}} = \frac{\langle \sigma v \rangle}{m_\chi} \Omega_\chi^2 (1+z)^6 \rho_{\mathrm{c}}^2$$

Clustered $$\left(\frac{\mathrm{d}E}{\mathrm{d}V\mathrm{d}t}\right)_{\mathrm{INJ}}^{\mathrm{ann,boosted}} = [1+B(z)] \left(\frac{\mathrm{d}E}{\mathrm{d}V\mathrm{d}t}\right)_{\mathrm{INJ}}^{\mathrm{ann}}$$

Boost Factor

$$\left(\frac{\mathrm{d}E}{\mathrm{d}V\mathrm{d}t}\right)^{\mathrm{Boosted}}_{\mathrm{INJ}} = [1 + B(z)] \left(\frac{\mathrm{d}E}{\mathrm{d}V\mathrm{d}t}\right)^{\mathrm{HMG}}_{\mathrm{INJ}}$$

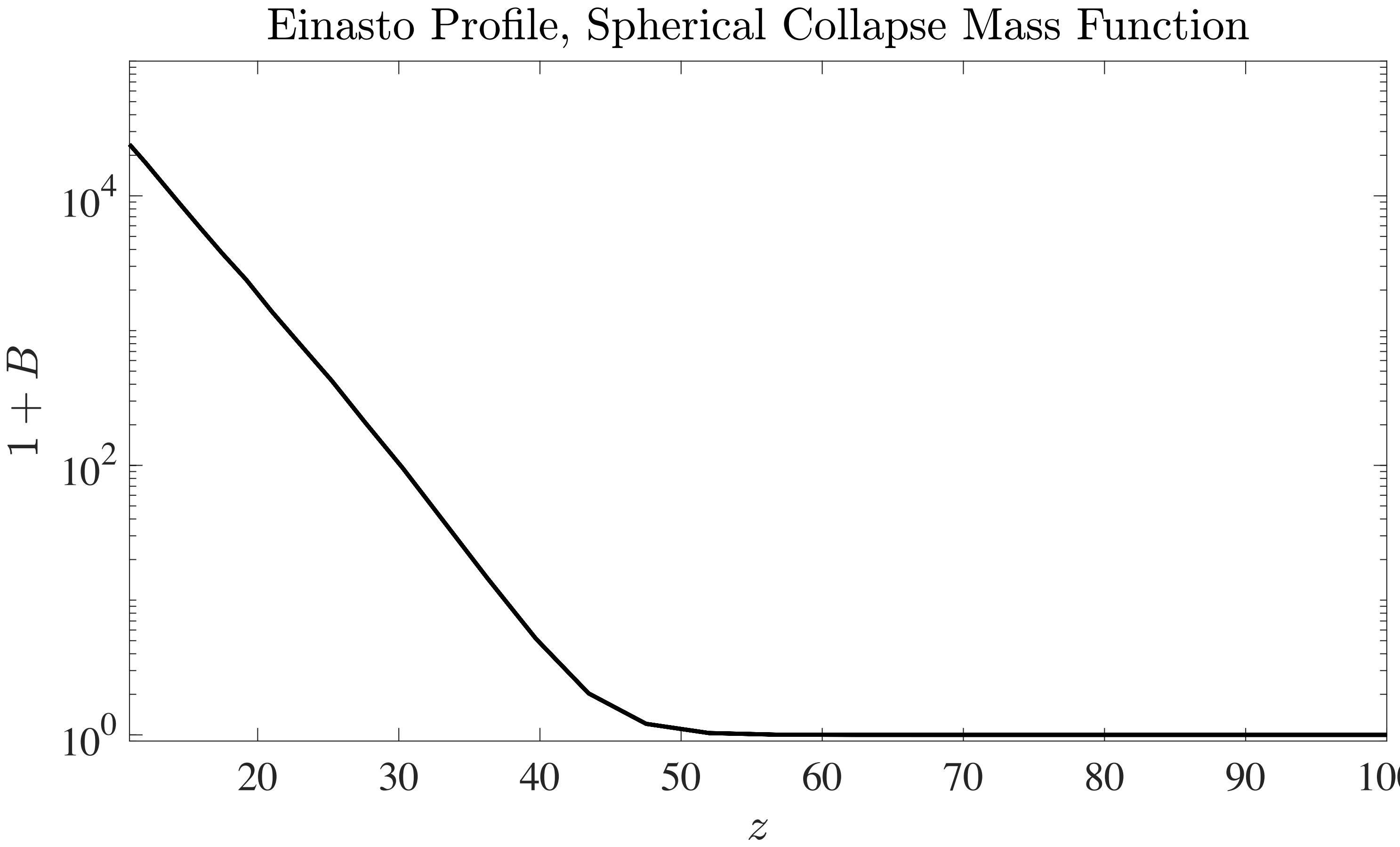
Boost Factor

$$B_{\mathrm{h}}(M) = \frac{4\pi}{\bar{\rho}_{\mathrm{h}}^2 V_{\mathrm{h}}(M)} \int_0^{r_{200}} \mathrm{d}r \rho^2(r) r^2$$

Halo Profile

$$B(z) = \frac{\Delta_c \rho_c}{\rho_{\mathrm{DM}}^2} \int_{M_{\mathrm{min}}}^{\infty} M B_{\mathrm{h}}(M) \frac{\mathrm{d}n}{\mathrm{d}M} \mathrm{d}M$$

Halo Mass Function



$$\rho(r) = \rho_{-2} \exp\left(-\frac{2}{\alpha_{\mathrm{e}}}\left[\left(\frac{r}{r_{-2}}\right)^{\alpha_{\mathrm{e}}} - 1\right]\right)$$

$$\frac{\mathrm{d}n}{\mathrm{d}\ln M} = \frac{1}{2} f(\nu) \frac{\rho_{\mathrm{DM}}}{M} \frac{\mathrm{d}\ln(\nu)}{\mathrm{d}\ln M}$$

$$f(\nu) = A\sqrt{\frac{2q\nu}{\pi}}[1 + (q\nu)^{-p}]\, \mathrm{e}^{-q\nu/2}$$

PBH Injection : Hawking Radiation

Particle Spectra:

$$\frac{dN}{dEdt} = \frac{1}{2\pi} \frac{\Gamma_s}{\exp(E/T_{BH}) - (-1)^{2s}}$$

Temperature

$$T_{\rm BH} = 1.06\,{\rm TeV}\frac{10^{10}{\rm g}}{M}$$

Mass Loss:

$$\dot{M} = -5.34 \times 10^{25} F(M) M^{-2} {\rm g/s}$$

$$F(M) : 1 \sim 15$$

$$\left(\frac{{\rm d}E}{{\rm dV d}t}\right)_{\rm INJ}^{\rm MMD} = -\dot{M} \cdot n_{\rm PBH}$$

Monochromatic Distribution:

$$n_{\rm PBH} = f_{\rm PBH} \frac{\Omega_{\rm DM} \rho_{\rm cr} (1+z)^3}{M}$$

$$f_{\rm PBH} \equiv \Omega_{\rm PBH} / \Omega_{\rm DM}$$

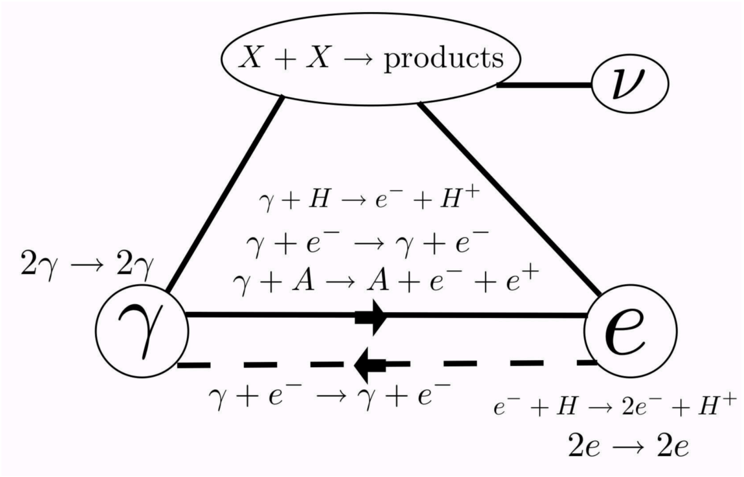
Extended Distribution:

$$\left(\frac{{\rm d}E}{{\rm dV d}t}\right)_{\rm INJ} = \int dM \Psi(M) \left(\frac{{\rm d}E}{{\rm dV d}t}\right)_{\rm INJ}^{\rm MMD}$$

$$\Psi(M) \equiv \frac{1}{\rho_{\rm PBH}} \frac{{\rm d}\rho_{\rm PBH}(M)}{{\rm d}M}$$

Deposition Efficiencies

$$\left(\frac{dE}{dtdV}\right)_{c,\mathrm{dep}} = f_c(z)\left(\frac{dE}{dtdV}\right)_{\mathrm{inj}}$$



- EM cascade
- Pair Production
- Compton Scattering
- Inverse CS

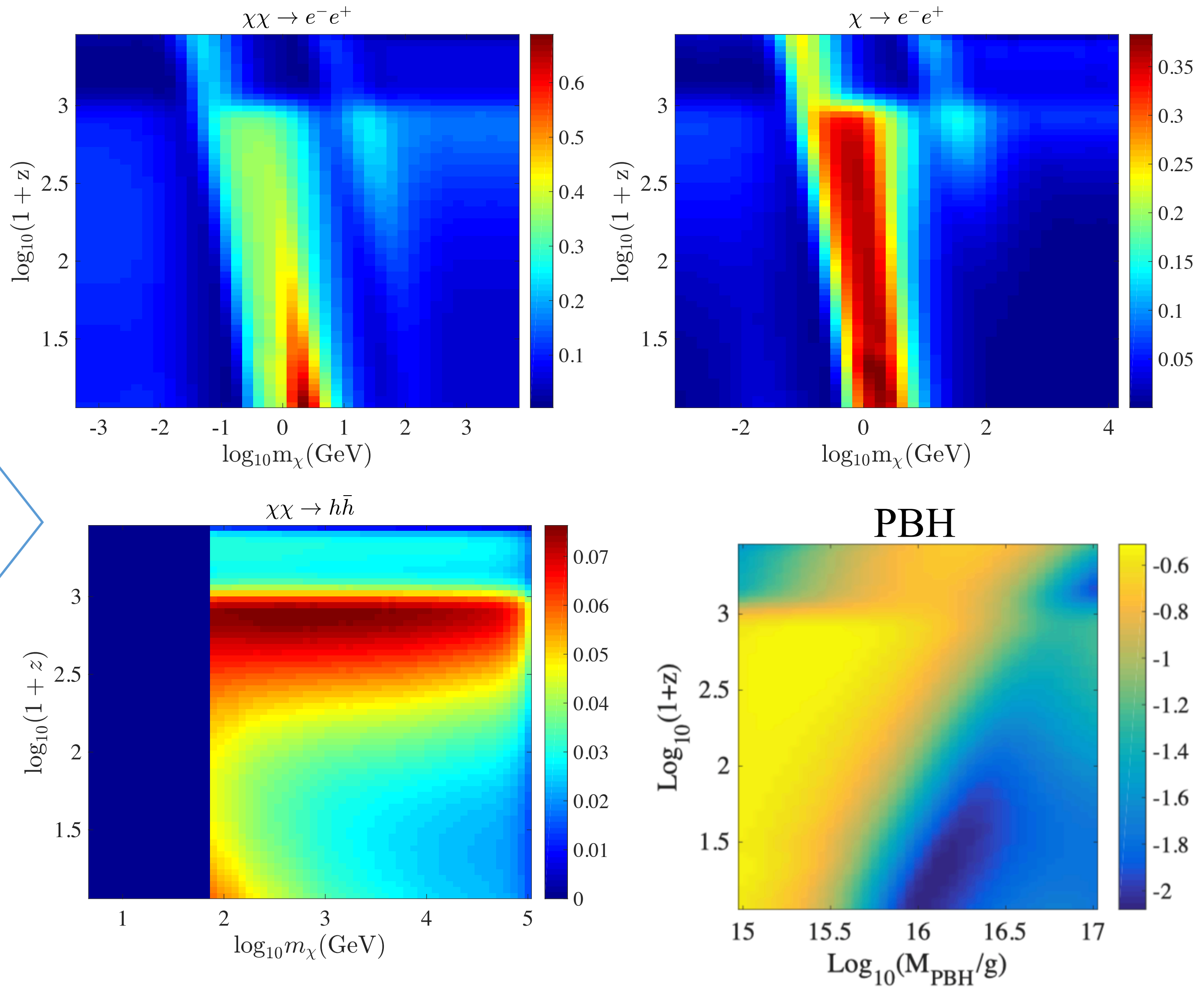
…

Dependencies

1. Redshift
2. Injection history
3. $e^{\pm}, \gamma$ Spectra
4. Particle Channel

……

Examples
(H Ionization)



Modified Recombination

$$\frac{\mathrm{d}x_\mathrm{e}}{\mathrm{d}z} = \left(\frac{\mathrm{d}x_\mathrm{e}}{\mathrm{d}z}\right)_0 - \frac{I_\chi}{(1+z)H(z)} \quad (16)$$

$$\frac{\mathrm{d}T_\mathrm{IGM}}{\mathrm{d}z} = \left(\frac{\mathrm{d}T_\mathrm{IGM}}{\mathrm{d}z}\right)_0 - \frac{2}{3k_B(1+z)H(z)}\frac{K_\mathrm{h}}{1+f_\mathrm{He}+x_\mathrm{e}}$$

From ground state (n=1)

Ionization Terms:

$I_\chi = I_{\chi i} + I_{\chi\alpha}$

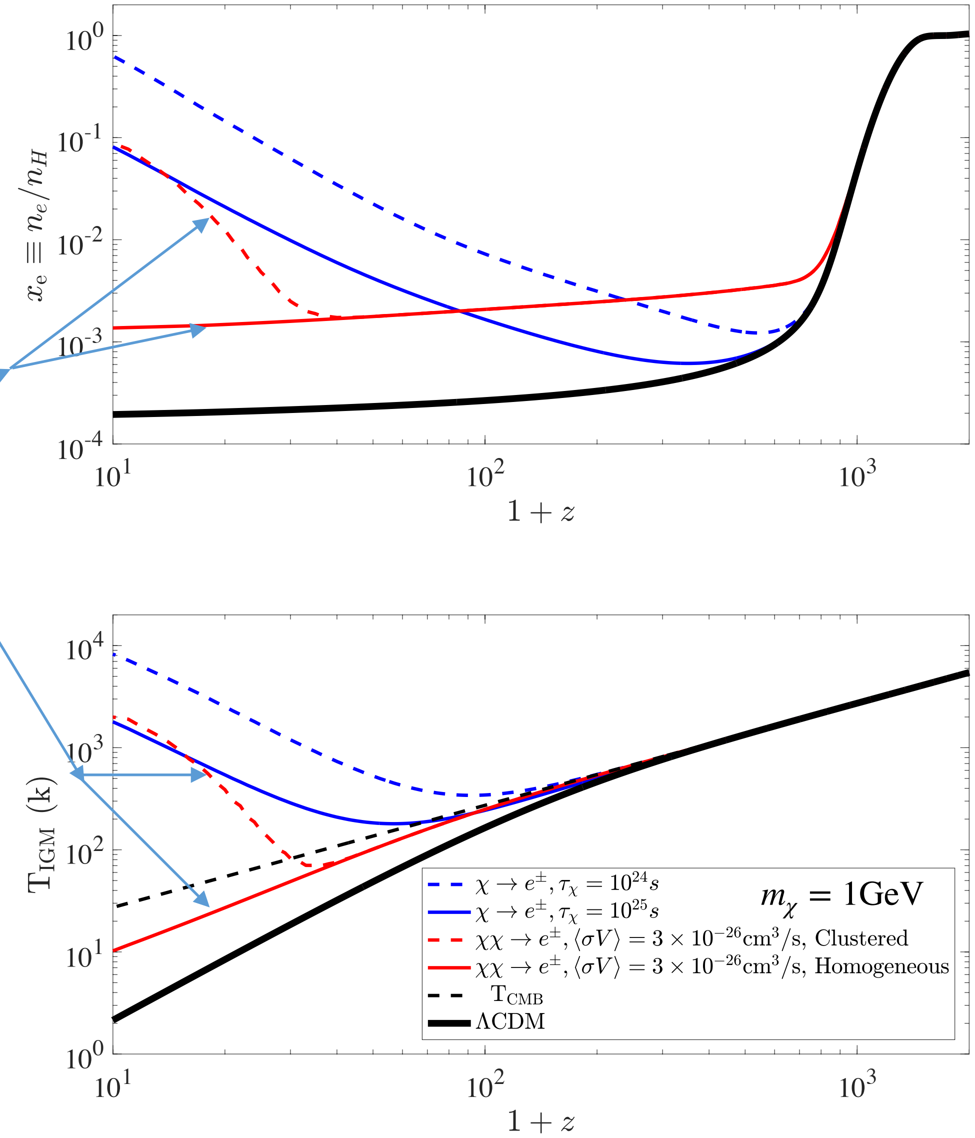
$$\begin{cases} I_{\chi i}(z) = \dfrac{1}{n_\mathrm{H}(z)E_i}\left(\dfrac{\mathrm{d}E}{\mathrm{d}V\mathrm{d}t}\right)_{\mathrm{DEP},i} \\ I_{\chi\alpha}(z) = \dfrac{1-C}{n_\mathrm{H}(z)E_\alpha}\left(\dfrac{\mathrm{d}E}{\mathrm{d}V\mathrm{d}t}\right)_{\mathrm{DEP},\alpha} \end{cases}$$

From excited state (n=2)

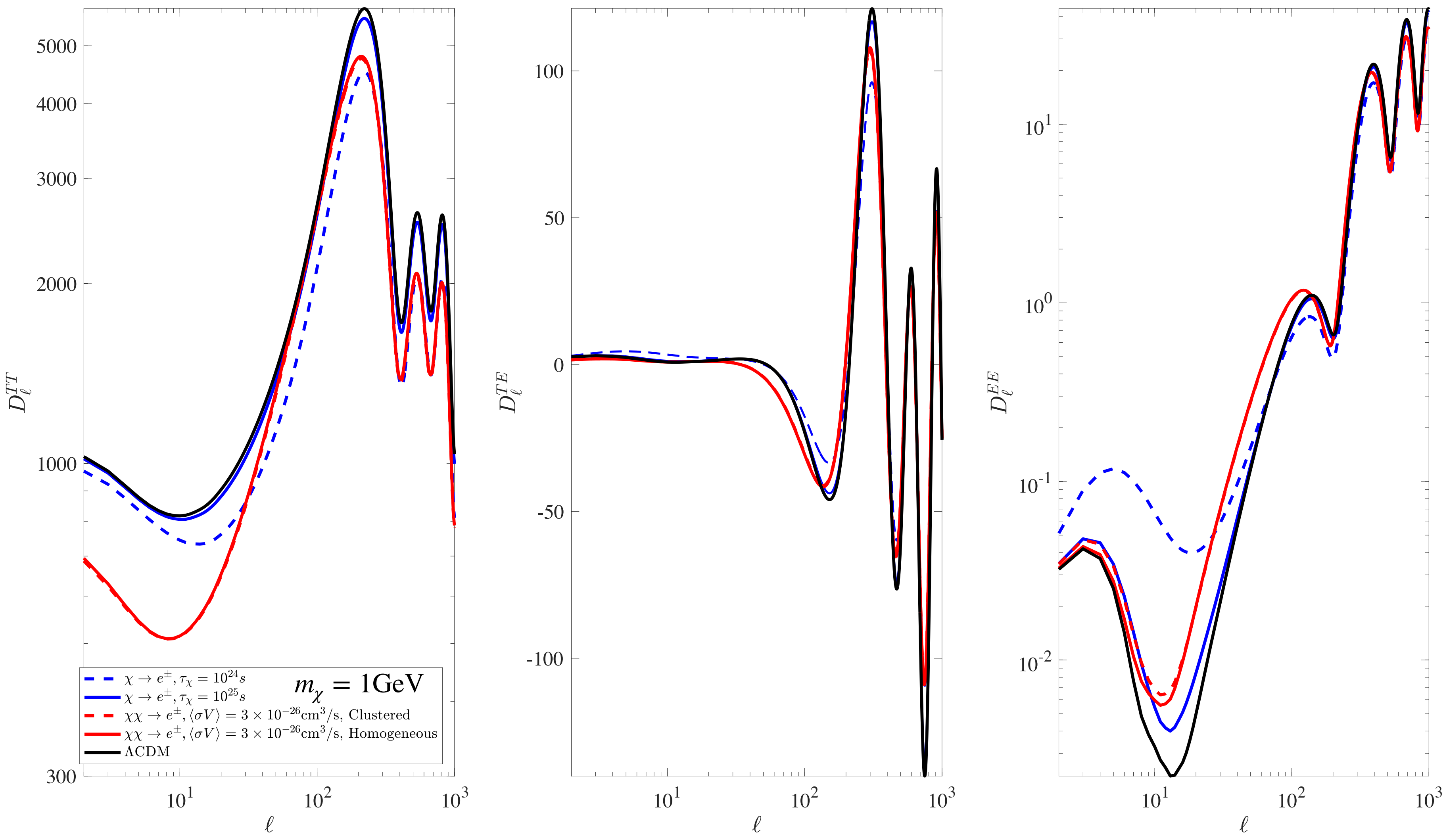
Heating Term:

$$K_\mathrm{h}(z) = \frac{1}{n_\mathrm{H}(z)}\left(\frac{\mathrm{d}E}{\mathrm{d}V\mathrm{d}t}\right)_{\mathrm{DEP},h}$$

Halo Boost



CMB Signature



Increase Compton Scattering

& Width of Surface of Last Scattering

1. Suppress small scale (high ℓ) correlations
2. Enhance Low-L correlations;
3. Shift EE and TE peak location

Insensitive to Halo Boost

Forecast Datasets

Chi^2

$$-2\ln\mathcal{L}(\{C_\ell\}|\{\hat{C}_\ell\}) = f_{\rm sky} \times \sum_\ell (2\ell+1)\{\mathrm{Tr}[\hat{C}_\ell C_\ell^{-1}] - \ln|\hat{C}_\ell C_\ell^{-1}| - 2\}$$

Data

$$\hat{C}_\ell = \bar{C}_\ell + N_\ell$$

ΛCDM   Noise

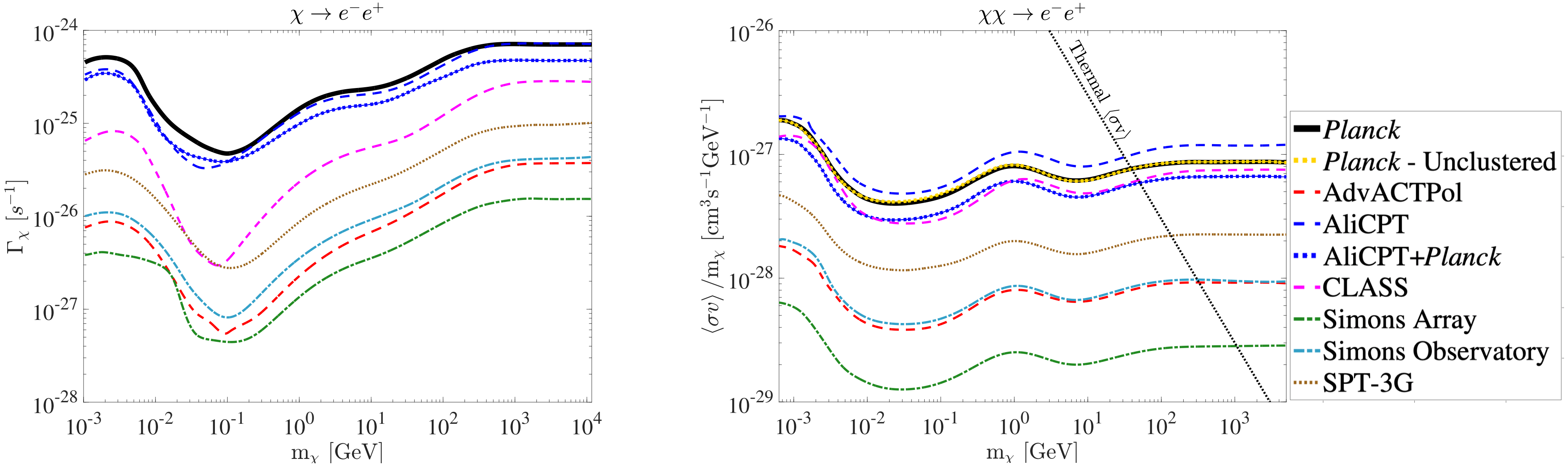
$$N_\ell^{\rm EE} = \left[\sum_\nu \frac{1}{N_{\ell,\nu}^{\rm EE}}\right]^{-1}, \quad N_\ell^{\rm TT} = N_\ell^{\rm EE}/2$$

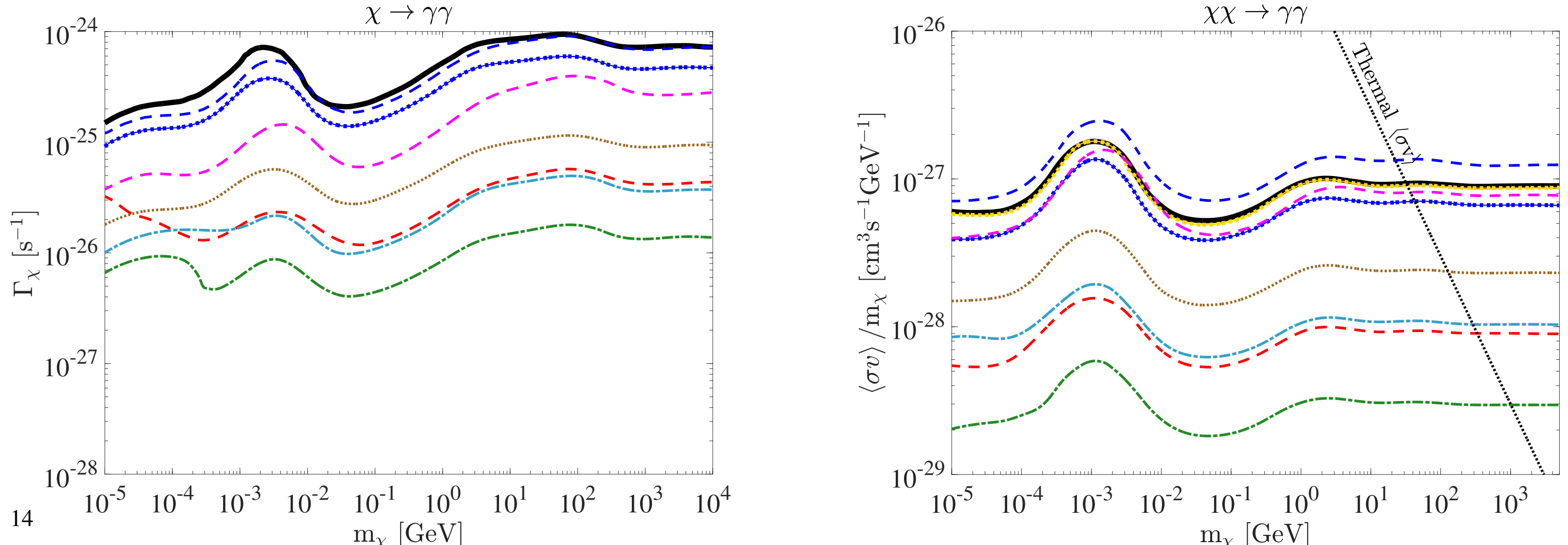
$$N_{\ell,\nu}^{\rm EE} = \delta P_\nu^2 \, \exp\left[\ell(\ell+1)\frac{\theta^2_{{\rm FWHM},\nu}}{8\,\ln 2}\right]$$

IHEP → AliCPT

Experiment	ν[GHz]	$\omega_{\rm E,\nu}^{-1/2}$ [μK-arcmin]	$\theta_{\rm FWHM}$[arcmin]	$f_{\rm sky}$[%]	$\ell_{\rm min}$	$\ell_{\rm max}$
AdvACTPol [20, 58, 59]	28 41 90 ⋆ 150 ⋆ 230	113.1 99.0 11.3 9.9 35.4	7.1 4.8 2.2 1.4 0.9	50	350[a]	4000
AliCPT [60]	90⋆ 150⋆	2 2	15.4 9.7	10	30	600
CLASS [22]	38 93⋆ 148⋆ 217	39 13 15 43	90 40 24 18	70	5	200
Simons Array [24, 61]	95⋆ 150⋆ 220	13.9 11.4 30.1	5.2 3.5 2.7	65	30	3000
Simons Observatory - SAT [25]	27 39 93⋆ 145⋆ 225 280	35.4 24 2.7 3 6 14.1	93 63 30 17 11 9	10	25	1000
Simons Observatory - LAT [25]	27 39 93⋆ 145⋆ 225 280	73.5 38.2 8.2 8.9 21.2 52.3	7.4 5.1 2.2 1.4 1 0.9	40	1000	5000
SPT-3G [19, 61, 62]	95⋆ 150⋆ 220	5.1 4.7 12.0	1 1 1	6	50	5000
CMB-S4 [70, 71]	95 145	2.9 2.8	2.2 1.4	62	30	3000
PICO [68, 69]	90 108 129 155	2.09 1.70 1.53 1.28	9.5 7.9 7.4 6.2	70	2	4000

DM Constraints (ArXiv 2002.03380)





Experiment	$\chi\chi \to e^+e^-$	$\chi\chi \to \gamma\gamma$
Planck	6.25	9.24
Planck - Unclustered	6.17	9.00
AdvACTPol	0.65	0.93
AliCPT	8.02	13.32
AliCPT+*Planck*	4.58	6.91
CLASS	4.85	8.26
Simons Array	0.20	0.31
Simons Observatory	0.68	1.08
SPT-3G	1.59	2.41

TABLE III. 95%C.L upper-bound on $\langle\sigma v\rangle / m_\chi$ (in $10^{-28} cm^3 s^{-1} GeV^{-1}$) for $m_\chi = 10$GeV.

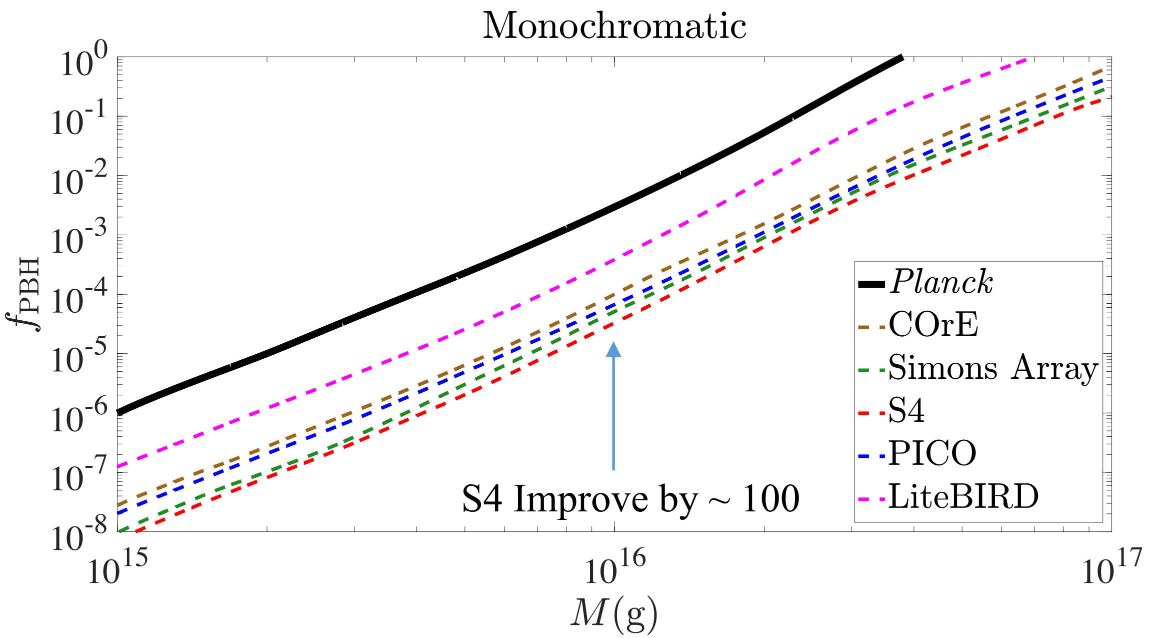
Experiment	$\chi \to e^+e^-$	$\chi \to \gamma\gamma$
Planck	23.61	85.02
AdvACTPol	0.68	4.66
AliCPT	20.77	78.31
AliCPT+*Planck*	16.00	53.07
CLASS	5.51	29.74
Simons Array	0.35	1.49
Simons Observatory	0.92	4.18
SPT-3G	2.20	9.88

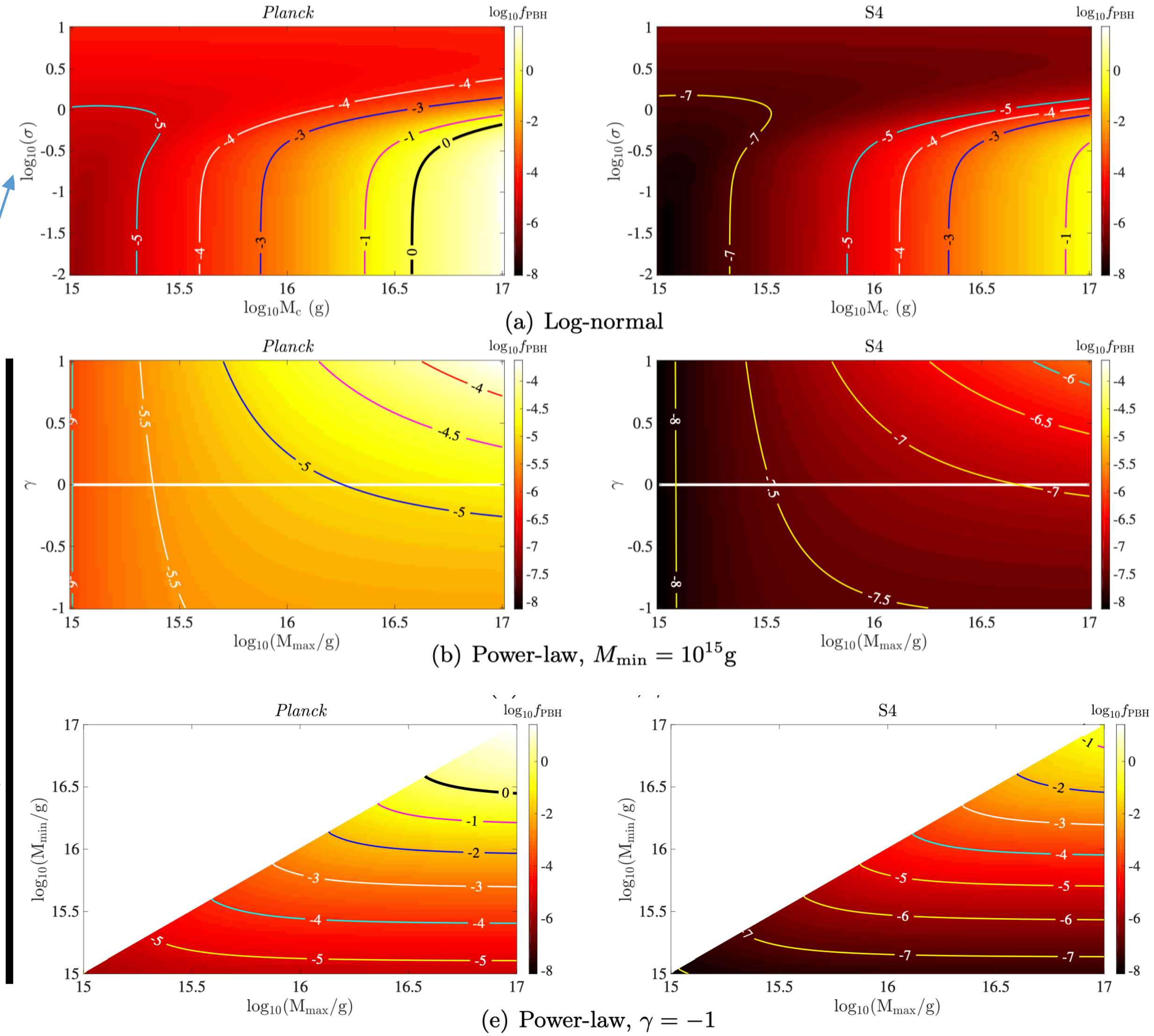
TABLE II. 95%C.L upper-bound on Γ_χ (in $10^{-26} s^{-1}$) for $m_\chi = 10$GeV.

- Strong, Robust Constraints (think about cosmic ray)
- Wide mass coverage : KeV (WDM) to TeVs (CDM)
- Prospective improvements by about a factor of 100
- **Insensitive to Halo Boost (Curse/Blessing?)**

21cm can solve this problem!

PBH Constraints (2011.12244)



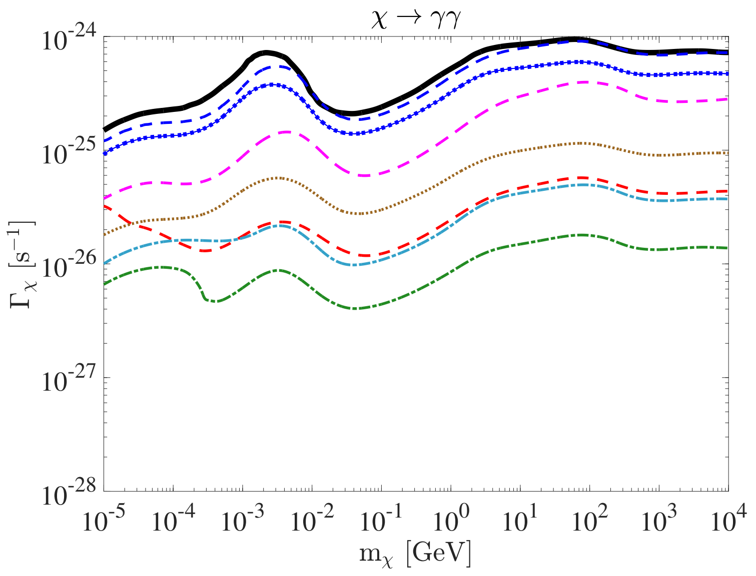


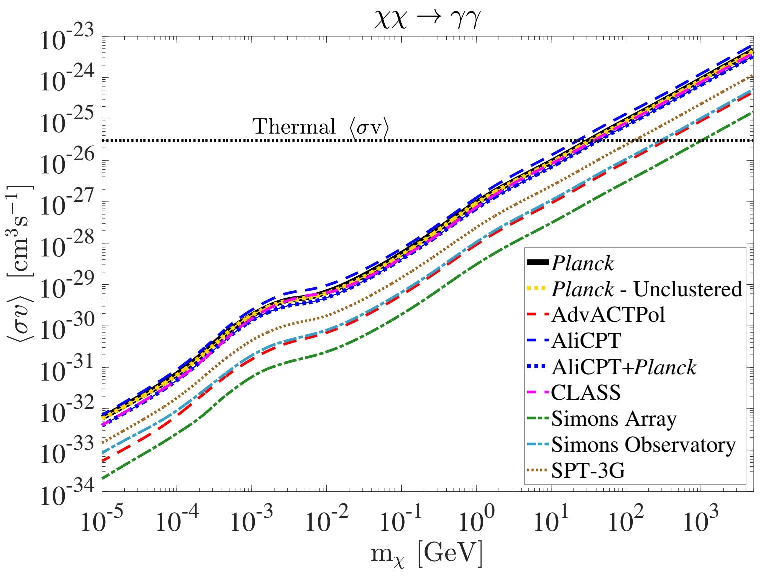
$$\Psi(M') = \frac{1}{\sqrt{2\pi}\sigma M'} \exp\left(-\frac{(\log[M'/M])^2}{2\sigma^2}\right)$$

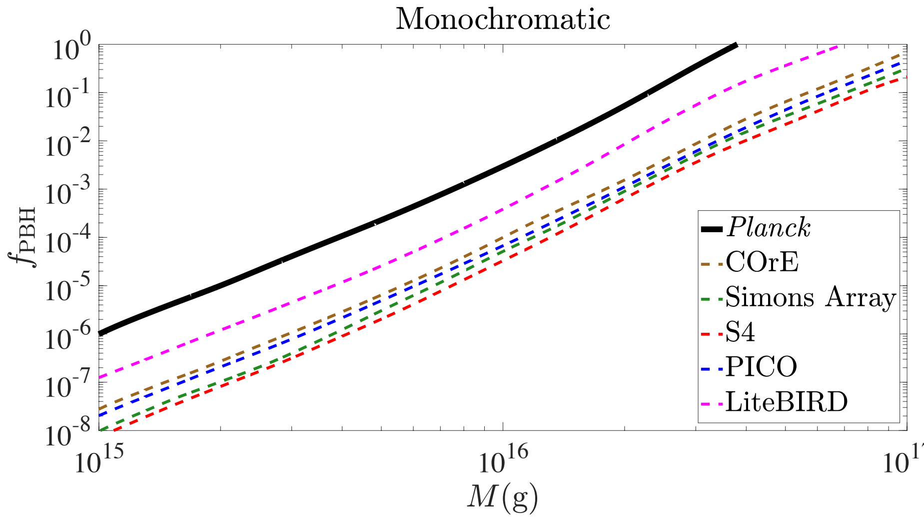
$$\Psi(M') \propto M'^{\gamma-1} \; (M_{\rm min} < M' < M_{\rm max})$$

$10^{15} \sim 10^{17}$g window can be ruled out

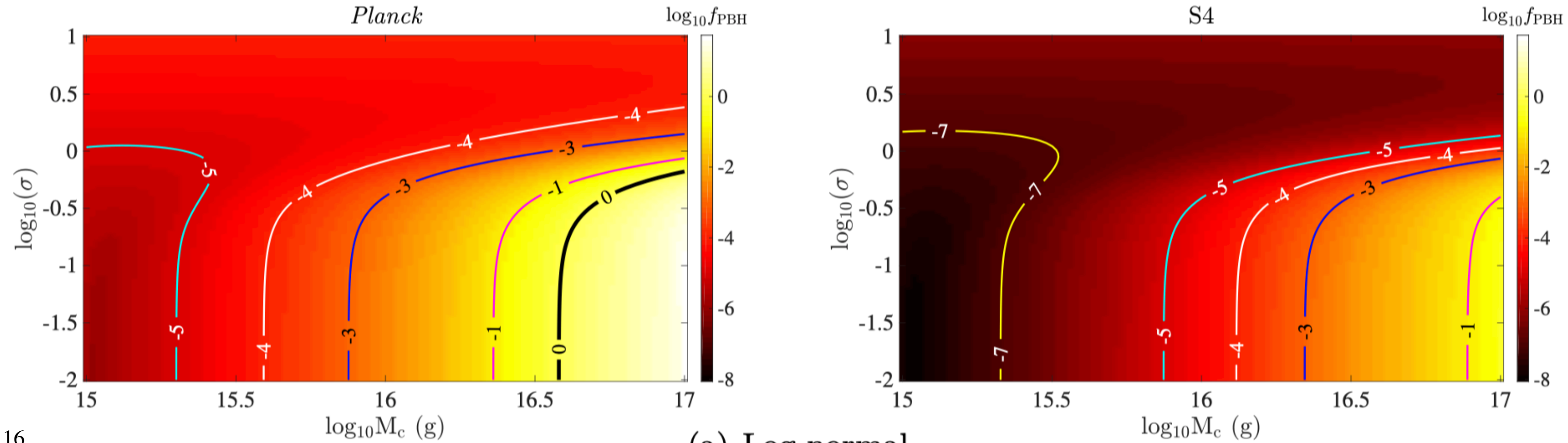
Thank you!







1. CMB can set robust and stringent DM/PBH bounds
2. Prospective improvement by 100 for future experiment
3. Stay tuned for the first 21cm light - SKA, HERA



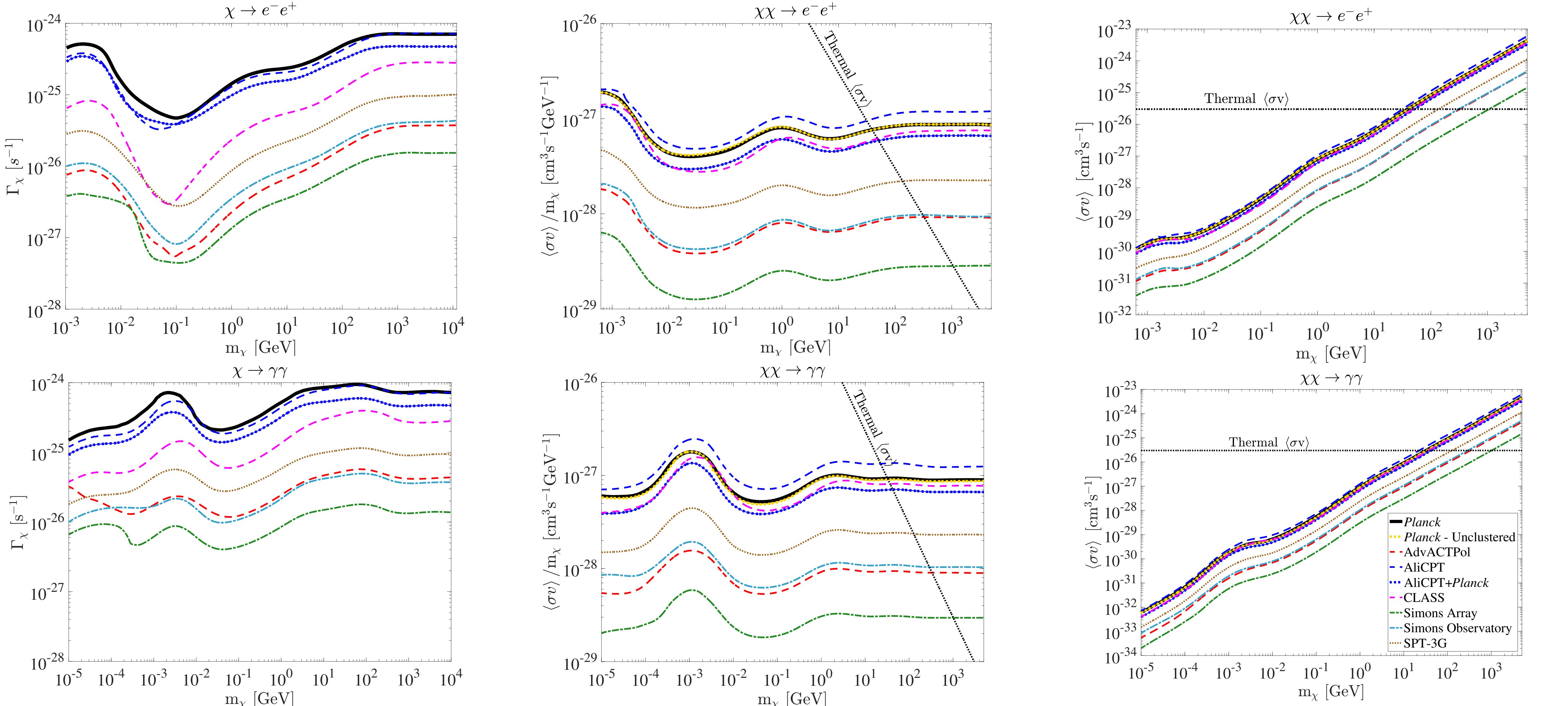
(a) Log-normal

Energy

Injection and Deposition

from DM

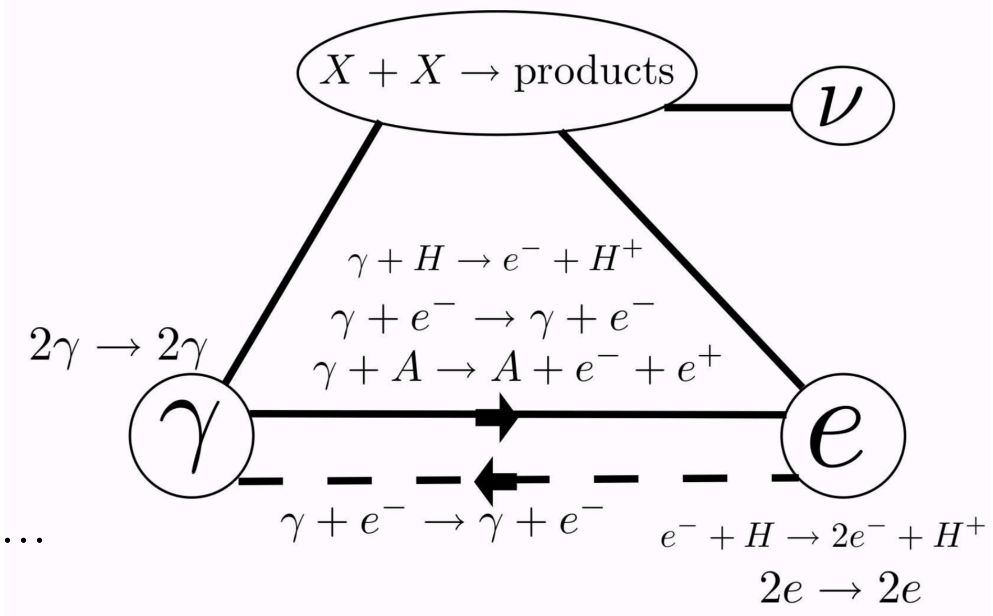
DM Constraints (ArXiv 2002.03380)



IGM Interaction

Most SM particles are either unstable or inactive.

Only $e^{\pm}, \gamma$ need to be accounted for.



$$I^s \equiv \frac{\mathrm{d}N^s}{\mathrm{d}E\mathrm{d}V\mathrm{d}t}$$

- DM: $I^s = \frac{\mathrm{d}N^s}{\mathrm{d}E} \times R$
  - $\frac{\mathrm{d}N^s}{\mathrm{d}E}$ ← Pythia, HERWIG, PPPC4DMID…
- PBH
  - Monochromatic

    $$I^{s,\delta}(M,t) = \frac{dN^s}{d\epsilon dt}(M,t) \times \frac{f_{\mathrm{BH}}\rho_c(z)}{M'}$$

    $$\frac{dN^s}{d\epsilon dt}(M,\epsilon) = \sum_i \int d\epsilon_i \frac{dN^s}{d\epsilon}(\epsilon,\epsilon_i)\frac{dN_i}{d\epsilon_i dt}(M,\epsilon_i)$$

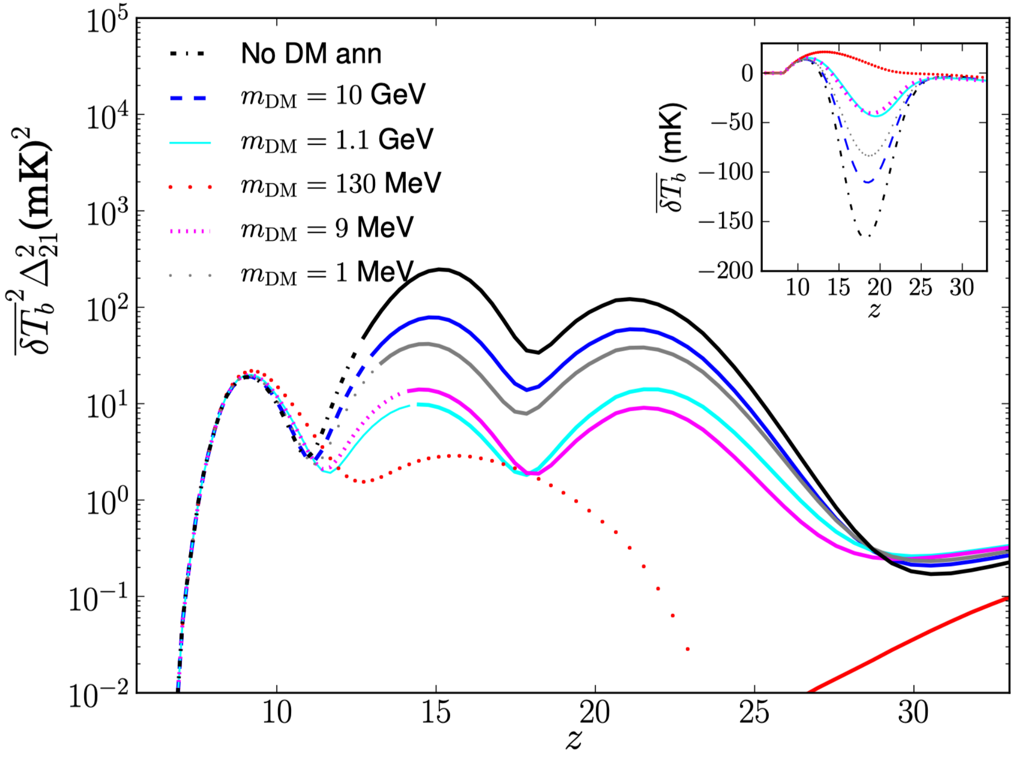
  - Extended

    $$I^s = \int_0^\infty dM \cdot \Psi(M) \cdot I^{s,\delta}(M,z)$$

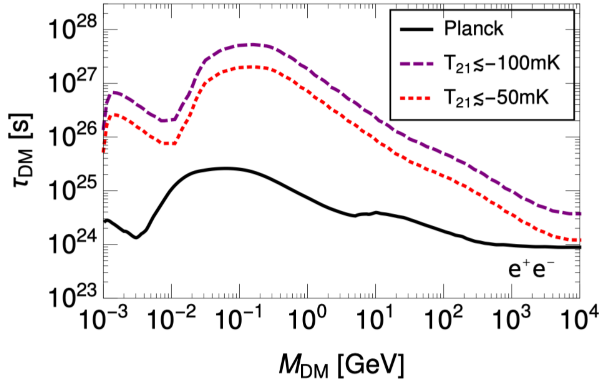
21cm Probe

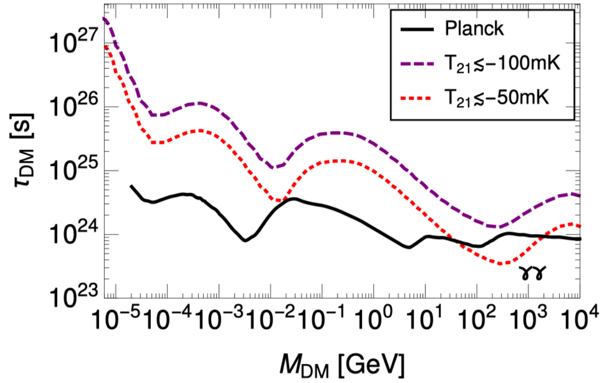
$$\delta T_b(\nu) \simeq 27\, x_{\rm HI}\, (1+\delta_b) \left(1 - \frac{T_{\rm CMB}}{T_S}\right) \left(\frac{1}{1+H^{-1}\partial v_r/\partial r}\right) \left(\frac{1+z}{10}\right)^{1/2} \left(\frac{0.15}{\Omega_m h^2}\right)^{1/2} \left(\frac{\Omega_b h^2}{0.023}\right) \text{mK}$$

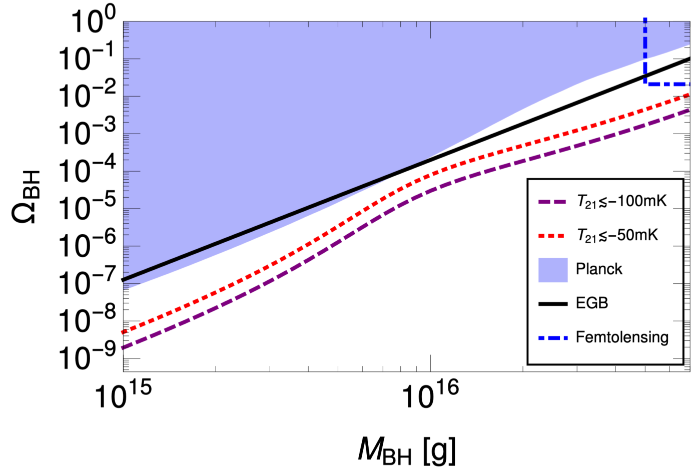
DM&PBH



ArXiv 1603.06795







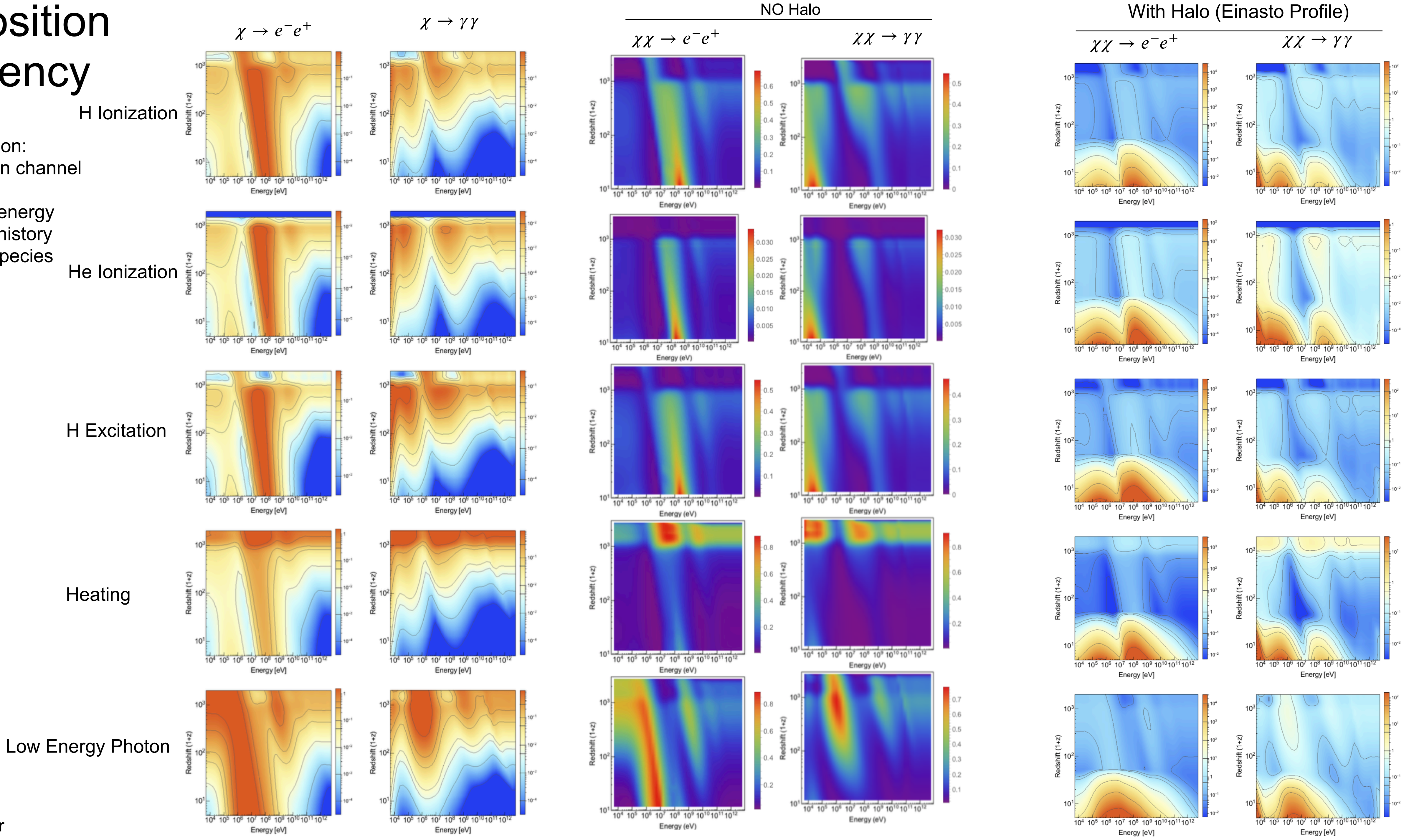
EDGES Constraints

ArXiv 1803.09390

Deposition Efficiency

Dependent on:

1. Deposition channel
2. Redshift
3. Injection energy
4. Injection history
5. Particle species

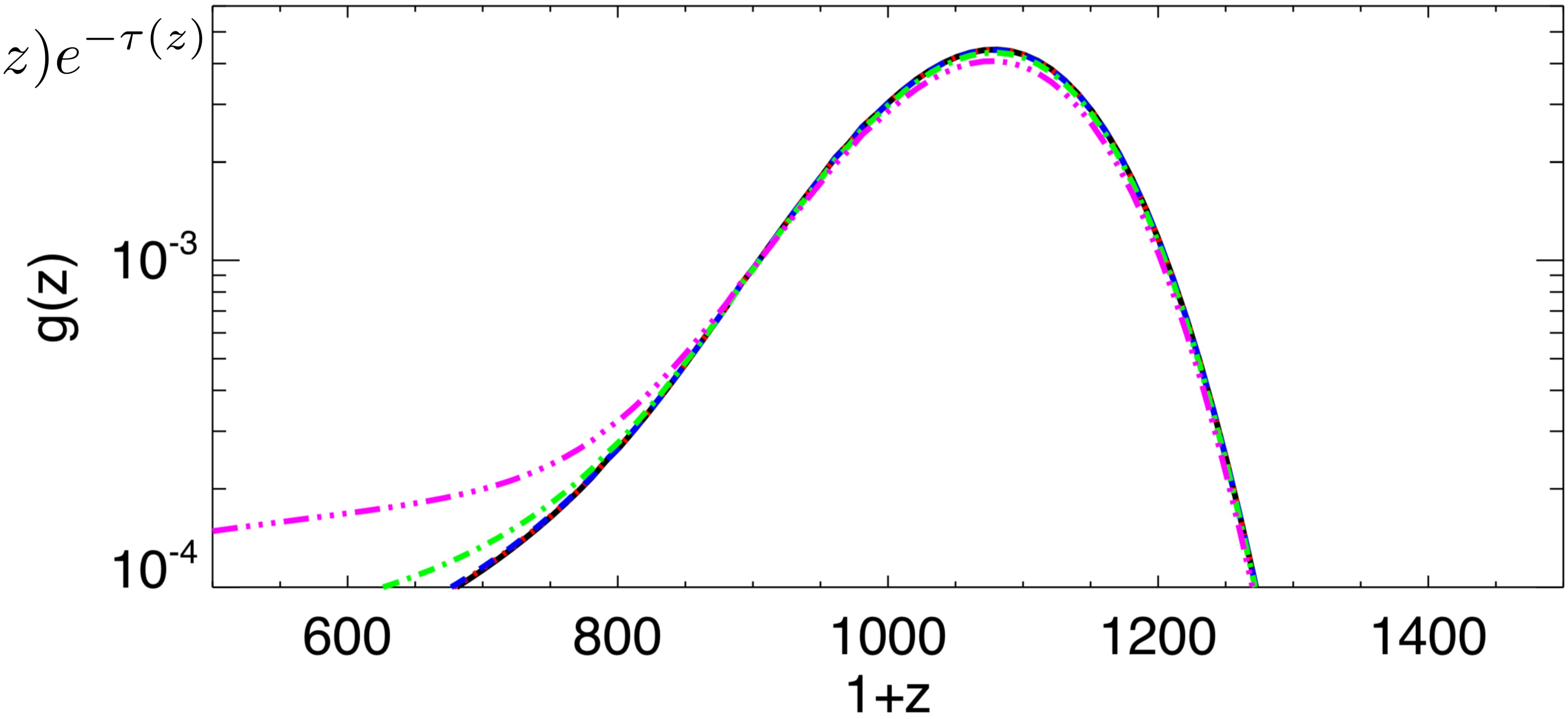


Courtesy
Of
Tracy Slatyer

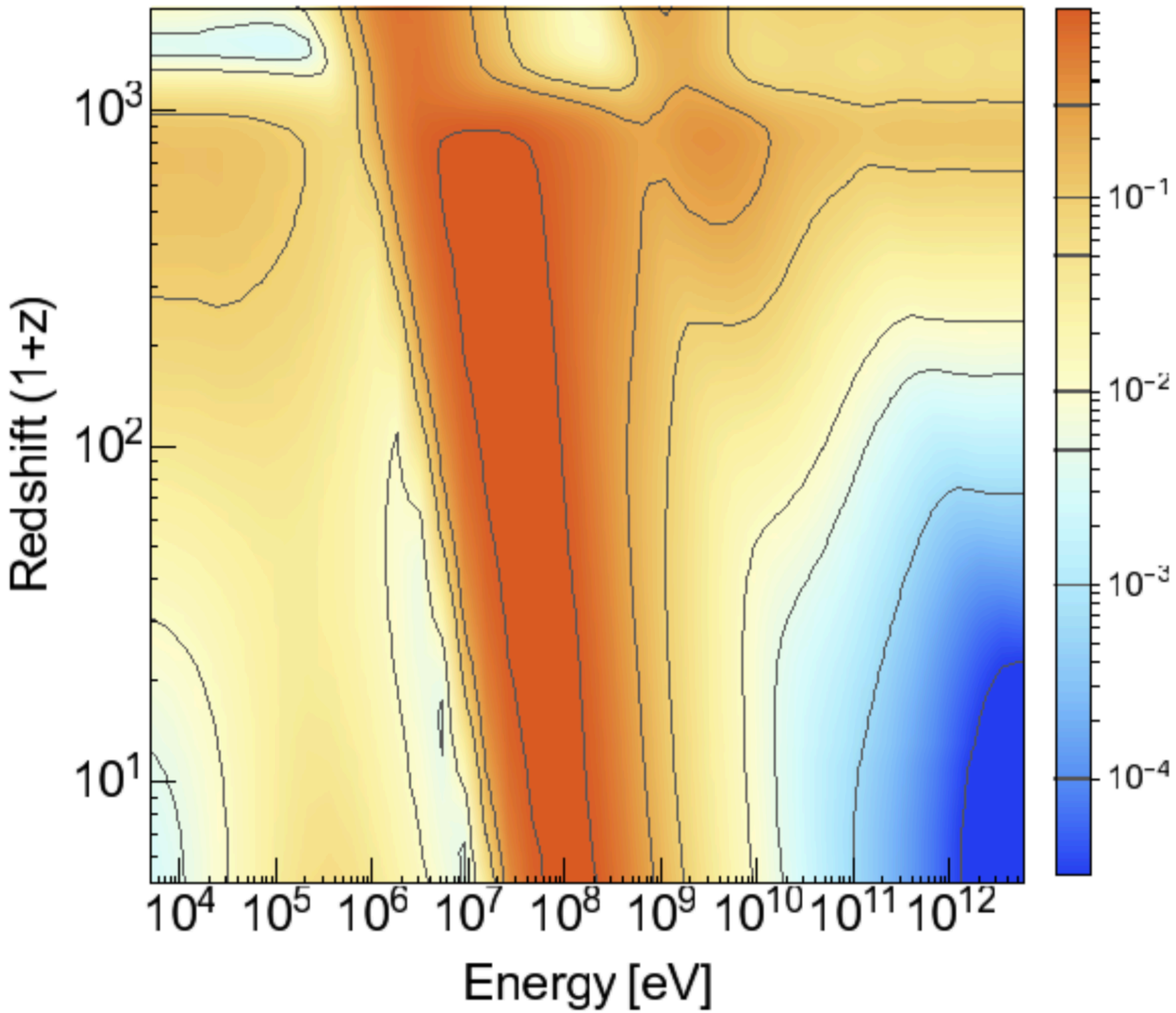
Width of laster scattering surface

Visibility function: probability that a photon last scattered between z and z + dz

$$g(z) \equiv \tau'(z)e^{-\tau(z)}$$



Stripe on e^-e^+Deposition Eff Figs



Location : 1 ~ 100 MeV

Electrons in this energy range upstater CMB photons to ~ 10 eV - KeV energies, where they can efficiently ionize hydrogen.