



Stockholm
University



PRINCETON
UNIVERSITY

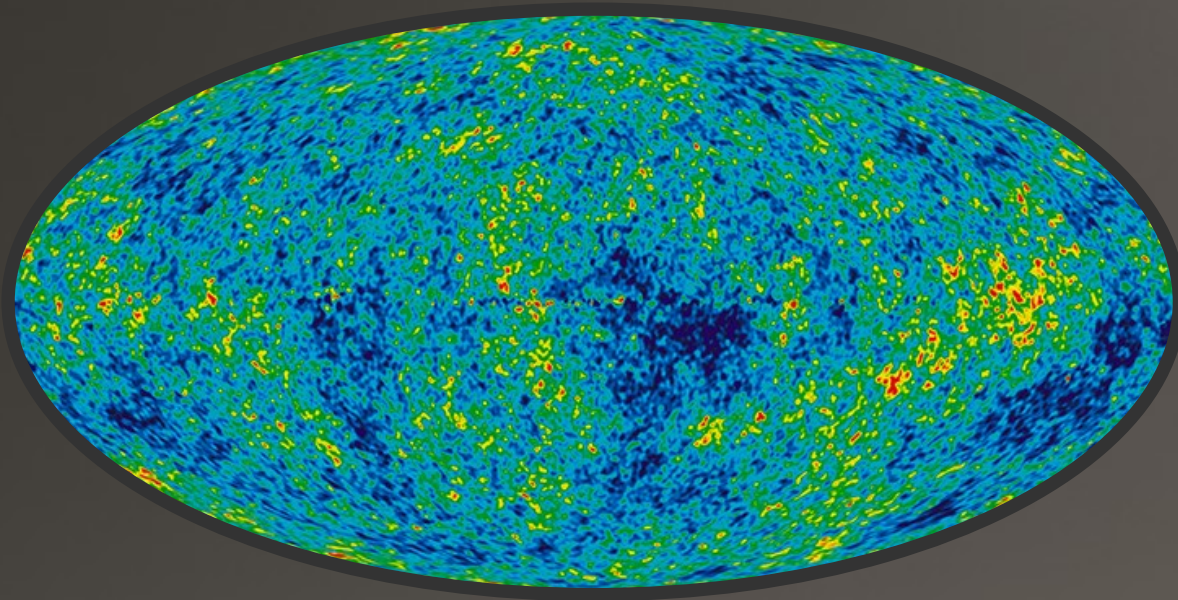
Developments in sub-GeV dark matter direct detection

CARLOS BLANCO

Outline

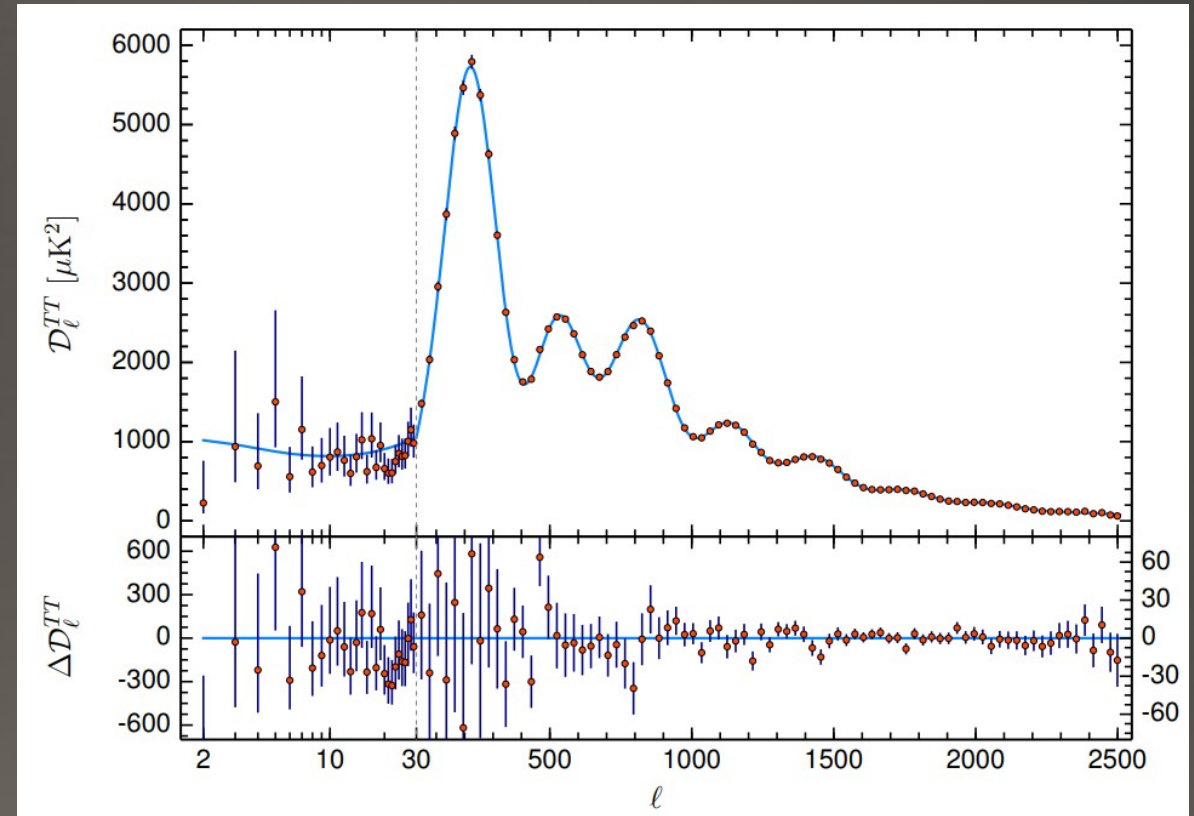
- Brief motivation for particle dark matter (DM)
- DM parameter space
- WIMPs – The miracle and tragedy
- Sub-GeV DM direct detection
 - Electron Recoil
 - Nuclear Recoil
- Directionality and modulating signals

Cosmological Evidence



ESA and the Planck Collaboration
CMB Anisotropies

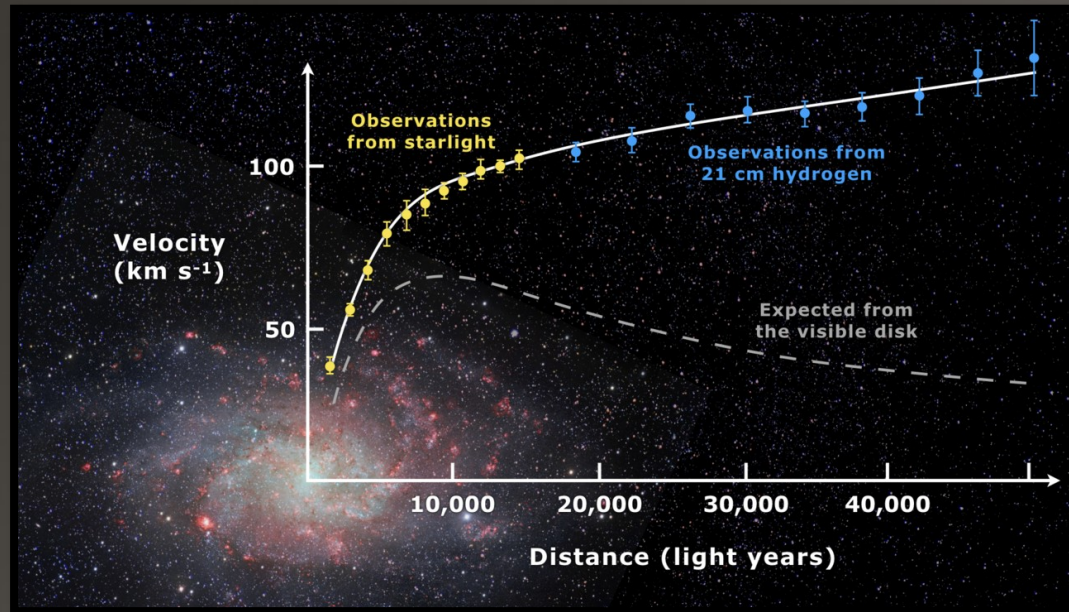
$$\Omega_c h^2 = 0.1198 \pm 0.0012$$



Planck results: 2020, A&A, 641, A6
2020, A&A, 641, A5

Astrophysical Evidence

Galactic rotation curves

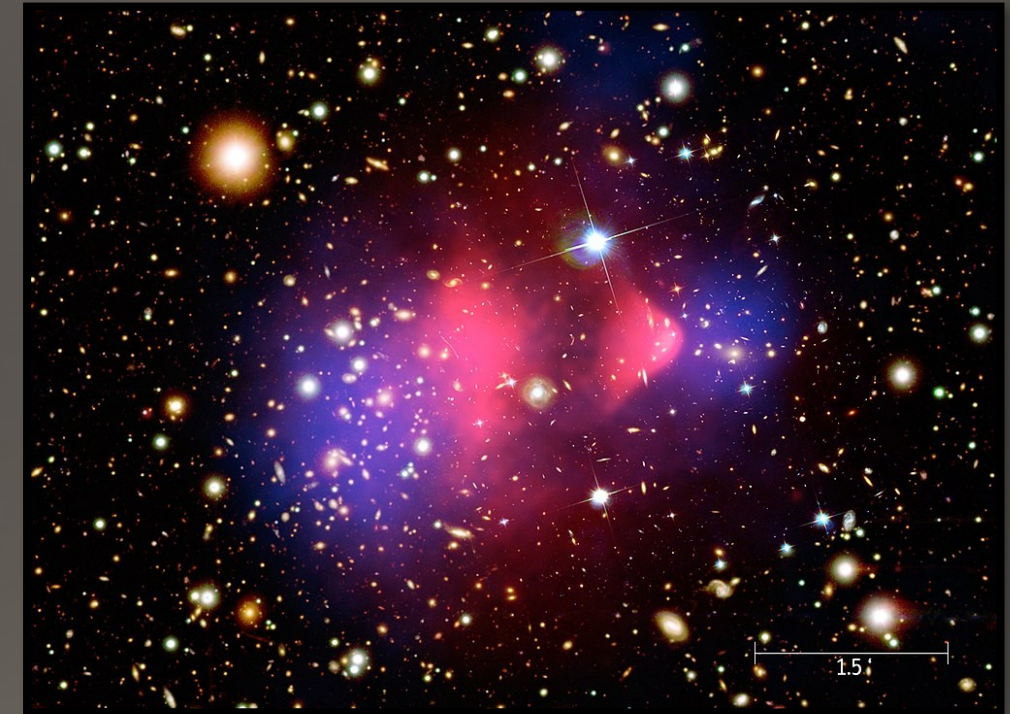


Solid: Observed

Dashed: Expected from visible disk

Corbelli & Salucci, MNRAS 1999
Bosma, Astronomical Journal, 1989

Bullet cluster

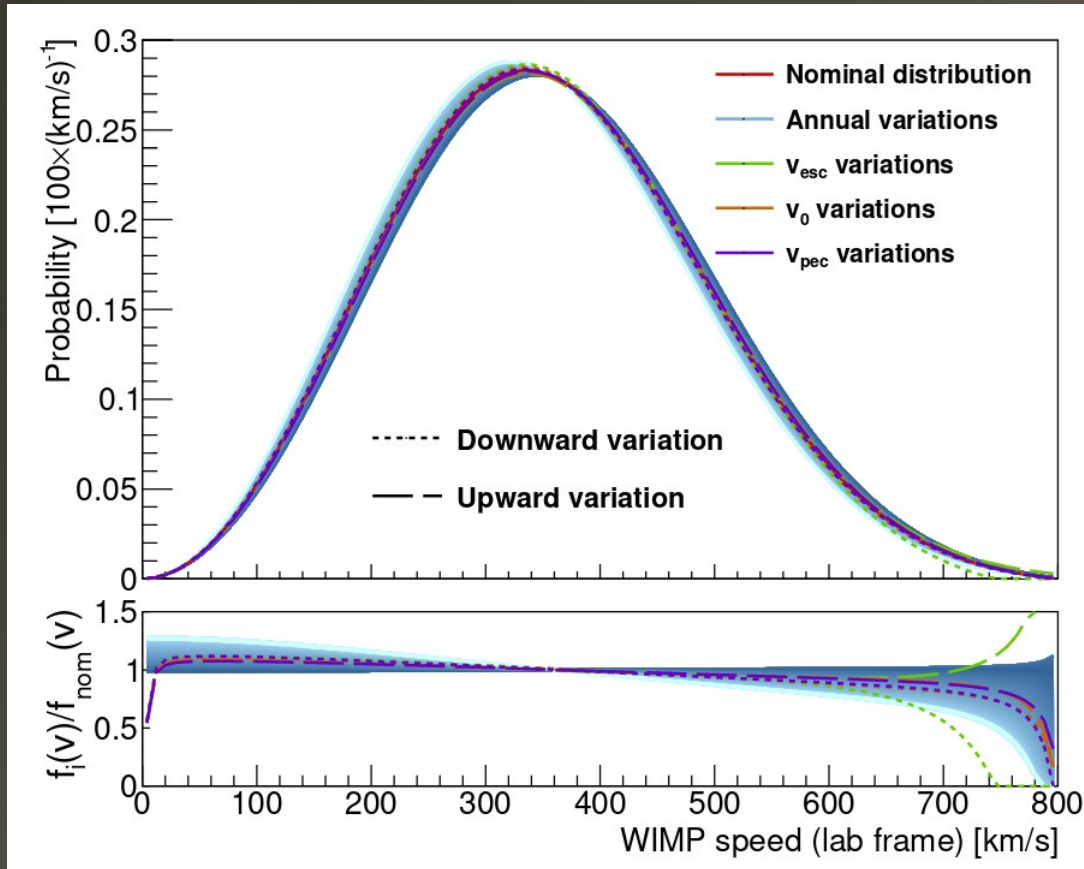


Blue: Mass from grav lensing

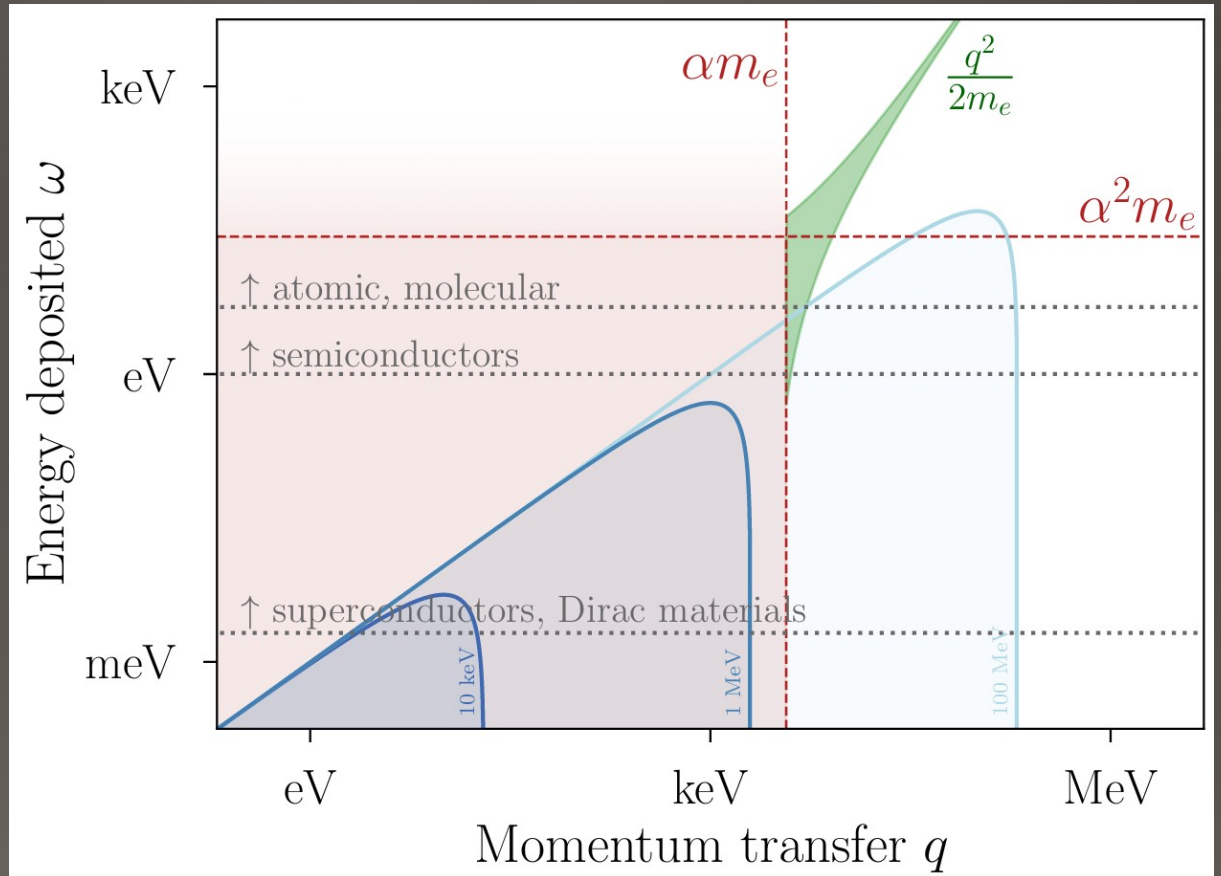
Pink: Baryonic mass from x-rays

X-ray: NASA/CXC/CfA/M.Markevitch et al.;
Optical: NASA/STScI; Magellan/U.Arizona/D.Clowe et al.;
Lensing Map: NASA/STScI; ESO WFI; Magellan/U.Arizona/D.Clowe et al.

Local DM Phase Space

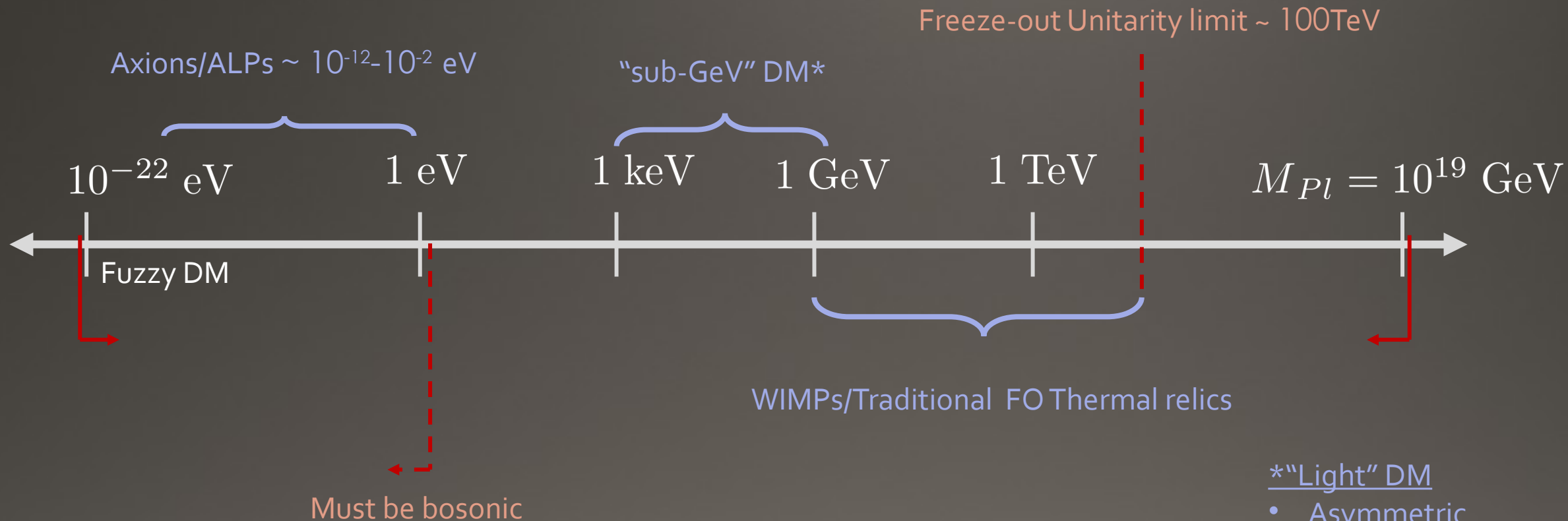


Baxter, D., et al. "Recommended conventions for reporting results from direct dark matter searches." *The European Physical Journal C* 81.10 (2021): 1-19.



Lin, Tongyan. "Sub-GeV dark matter models and direct detection." *SciPost Physics Lecture Notes* (2022): 043.

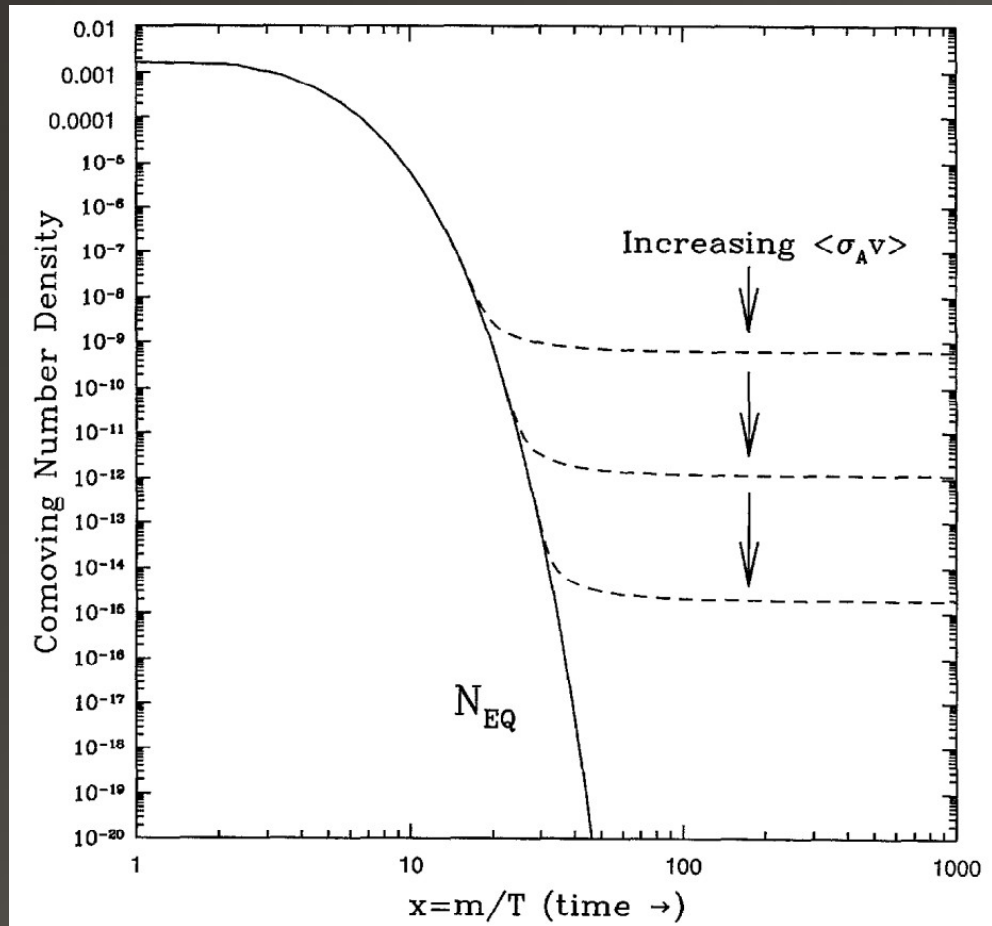
DM Mass Range



(NOT TO SCALE)

- *"Light" DM
- Asymmetric DM
- Hidden sector
- ELDER/SIMPs
- Freeze-In
- ...etc

WIMPs: The Miracle and Tragedy



Griest et. al: Phys Rep. 1996

$$\Omega_c h^2 \sim \frac{1}{\langle\sigma v\rangle}$$

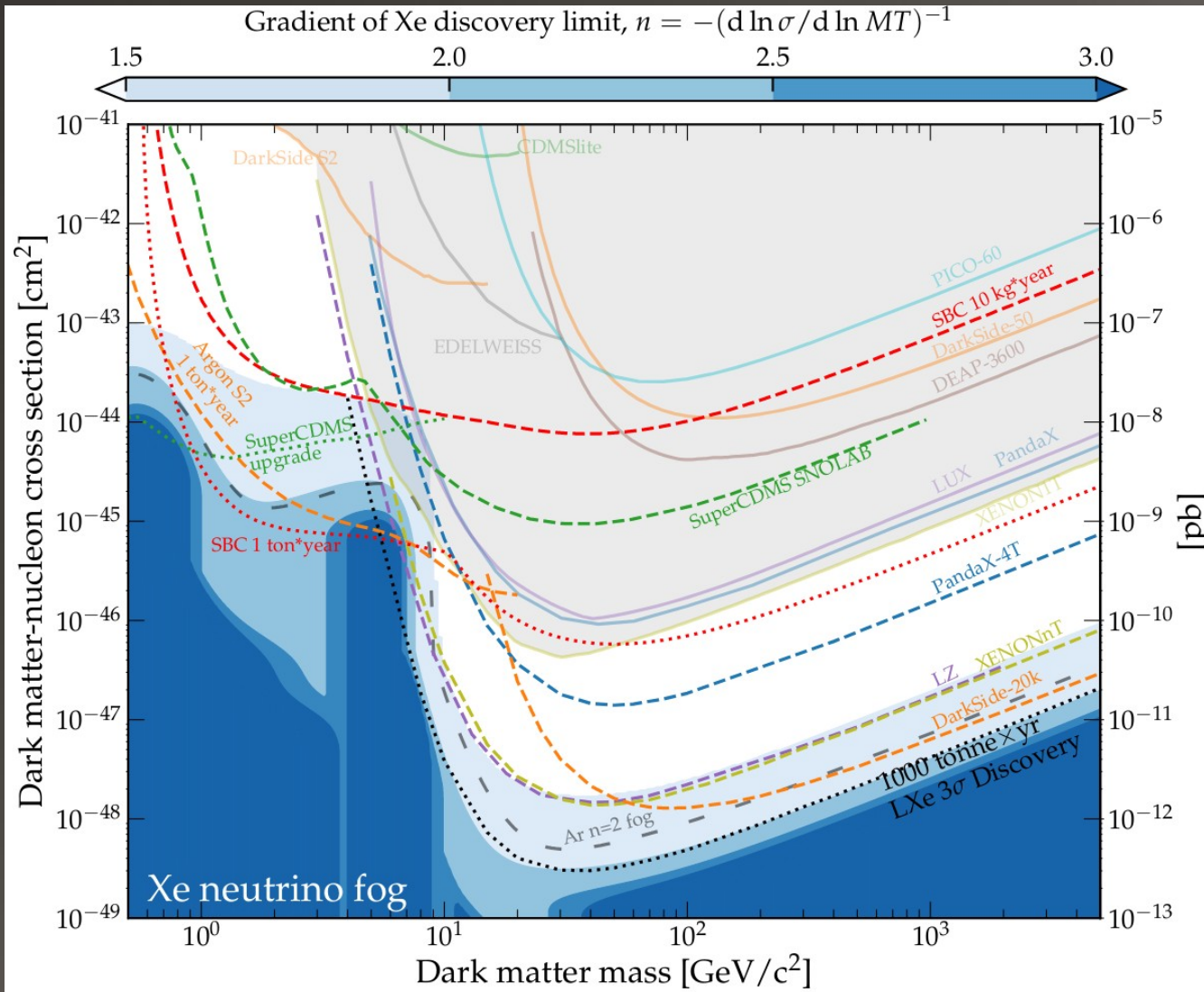
Predicts:

$$\langle\sigma v\rangle_{FO} \approx 3 \times 10^{-26} \text{ cm}^3 \text{ s}^{-1}$$

Subject to constraints from :

- CMB (recombination)
- Fermi-LAT (Dwarf spheroidal galaxy observations)
- BBN
- LEP, LHC, AMS,...etc

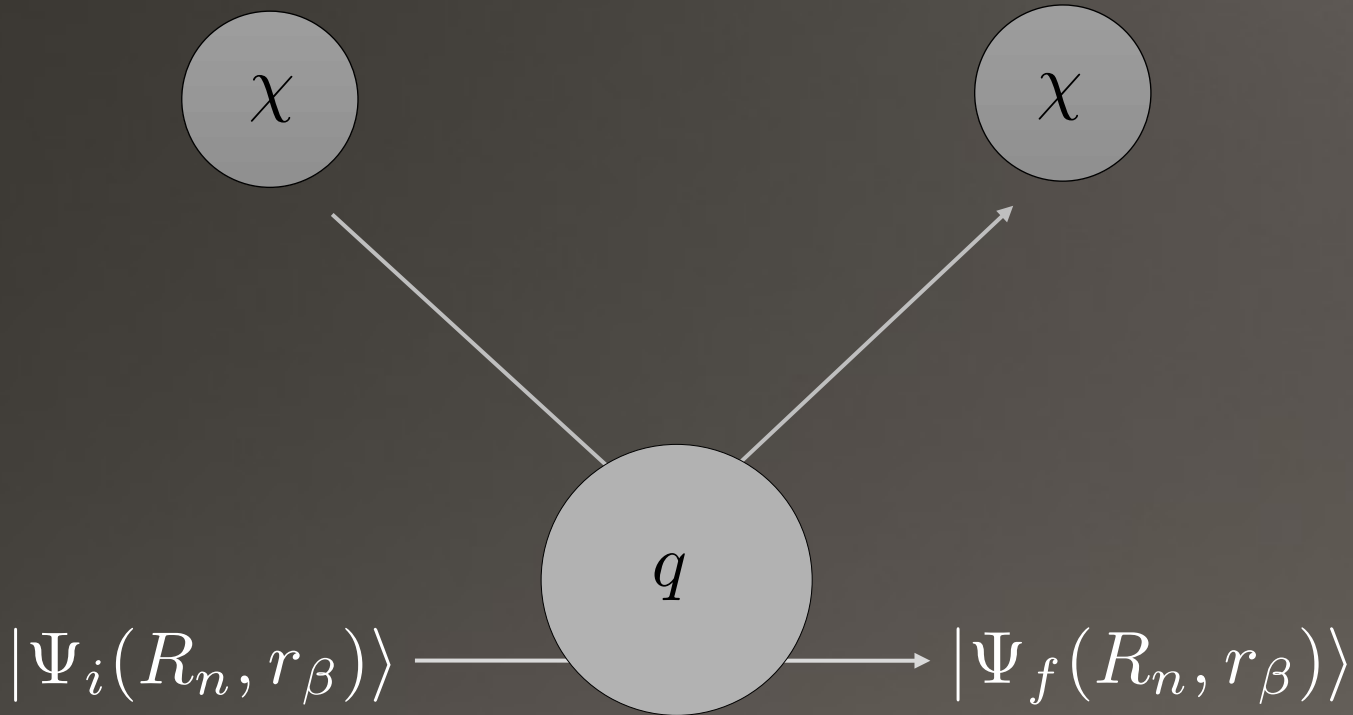
WIMPs: The Miracle and Tragedy



- Noble gasses: Very effective at probing masses above 1 GeV
- Will encounter irreducible neutrino fog in next generation
- Linear at high masses due to number density
- Loss of sensitivity due to threshold at low masses

Sub-GeV Direct Detection

$$\langle v_\chi \rangle \approx 300 \text{ km/s}$$



Kinematics

$$\Delta E = \vec{q} \cdot \vec{v}_\chi - \frac{q^2}{2\mu_{\chi,m}}$$

$$|q| \sim m_\chi v_\chi \approx \mathcal{O}(\text{KeV}) \left(\frac{m_\chi}{1 \text{ MeV}} \right)$$

R_i : Nuclear coordinates

r_β : Electron coordinates

Interaction Rate

$$\frac{dR}{d \ln E_r} = \frac{N_T \rho_\chi \bar{\sigma}_e}{8\pi \chi \mu_{\chi, m_T}^2} \int \frac{d^3 \vec{q}}{q} \eta(v_{\min}) |F_{\text{DM}}(q)|^2 |f_{i \rightarrow f}(q)|^2$$

Photon Rate

Mean inverse velocity

$$\eta(v_{\min}) = \int \frac{d^3 \vec{v}}{v} g(v) \Theta(v - v_{\min})$$

Dark matter form factor

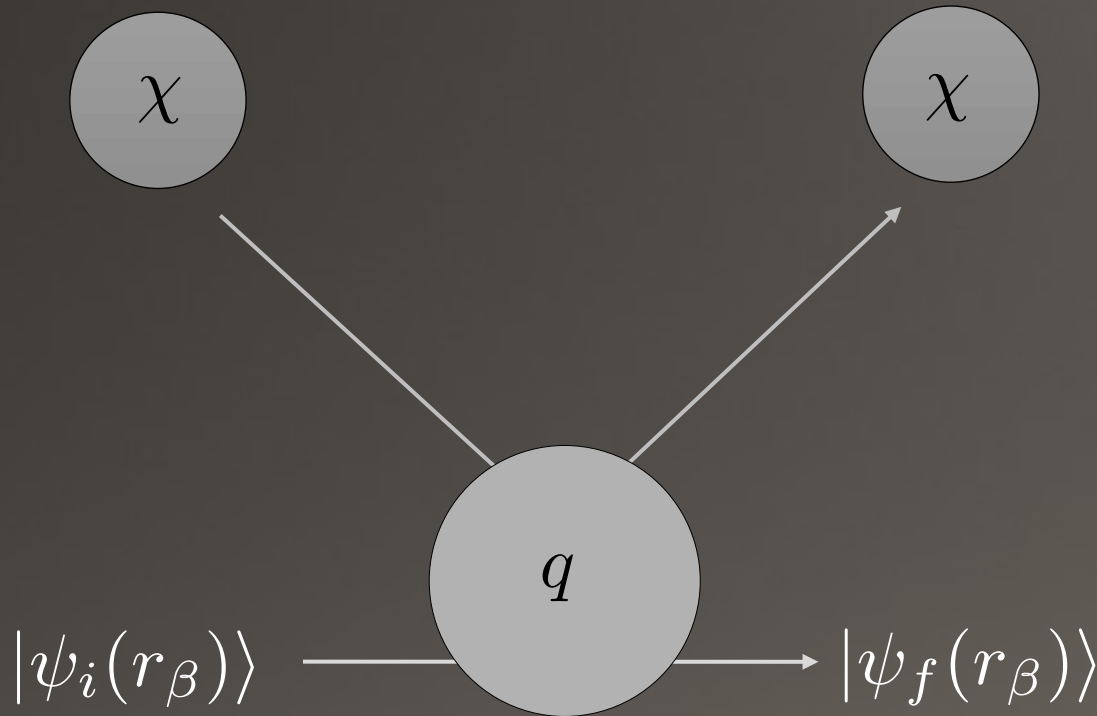
$$F_{\text{DM}}(q) = \begin{cases} 1 & , \text{ Heavy mediator} \\ \left(\frac{\alpha m_e}{q}\right)^2 & , \text{ Light mediator} \end{cases}$$

Transition form factor

$$f_{i \rightarrow f}(q) = \int d^3 \vec{R} d^3 \vec{r} \Psi_f^*(R_i, r_\beta) e^{i\vec{q} \cdot (\vec{r}_\beta, \vec{R}_i)} \Psi_i(R_i, r_\beta)$$

$$f_{i \rightarrow f}(q) = \int d^3 \vec{k} \tilde{\Psi}_f^*(k + q) \tilde{\Psi}_i(k)$$

Electron Recoil: Charge Signal



Electron scattering

$$\Delta E_r = (m_\chi^2 / m_T) \times 10^{-6}$$

$$\Delta E \sim \mathcal{O}(\text{few eV}) \left(\frac{m_\chi}{1 \text{ MeV}} \right)^2$$

What has such transition energies?

- Semiconductor band gaps
- Maybe atomic ionization

Electrons in crystals (exciton generation)

Electrons in atoms (ionization)

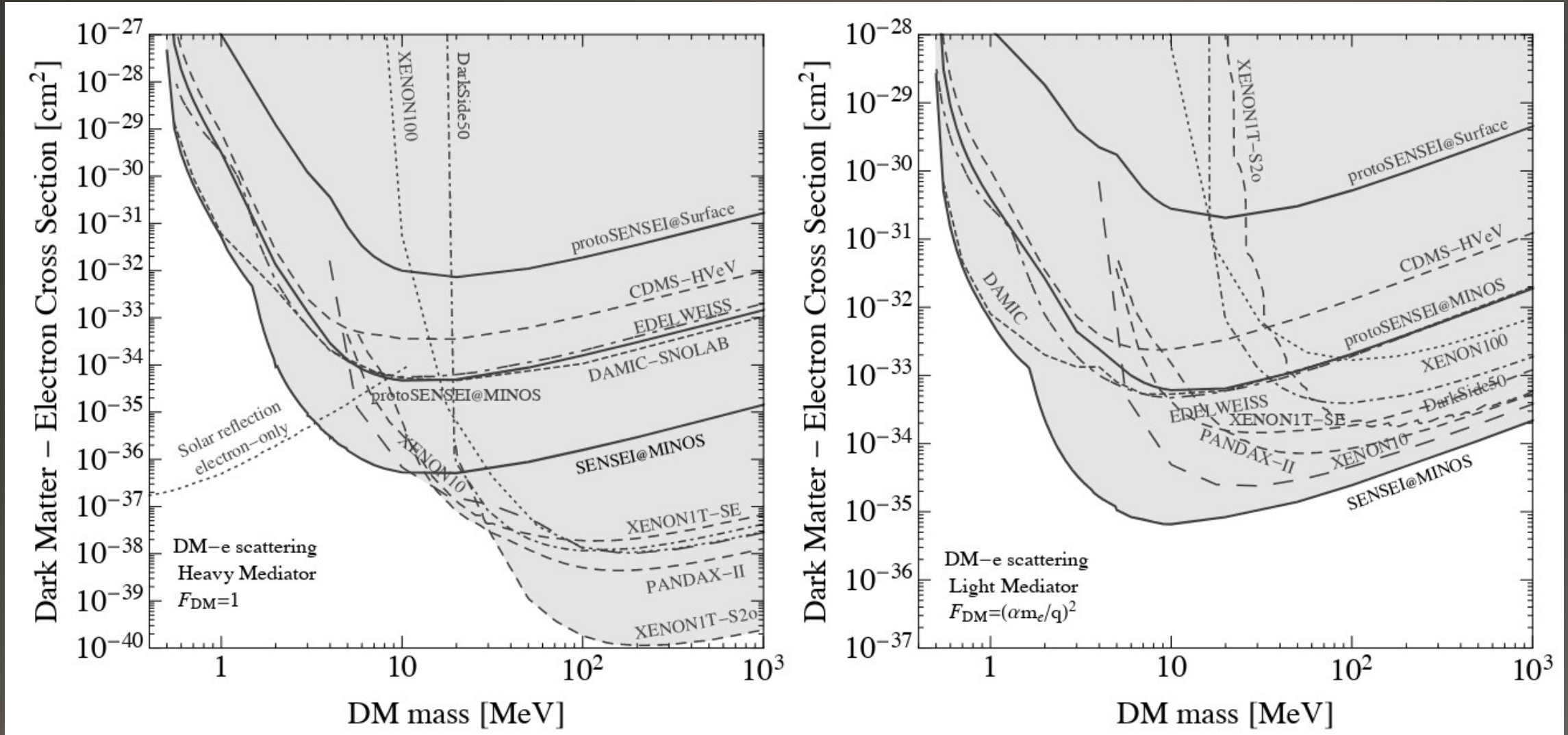
$$|\psi_i\rangle \sim u_v(r) e^{ik' \cdot r}$$

$$|\psi_f\rangle \sim u_c(r) e^{ik \cdot r}$$

$$|\psi_i\rangle \sim \psi_{\text{STO}}(r_\beta)$$

$$|\psi_f\rangle \sim e^{ik \cdot r}, r \gg a_0$$

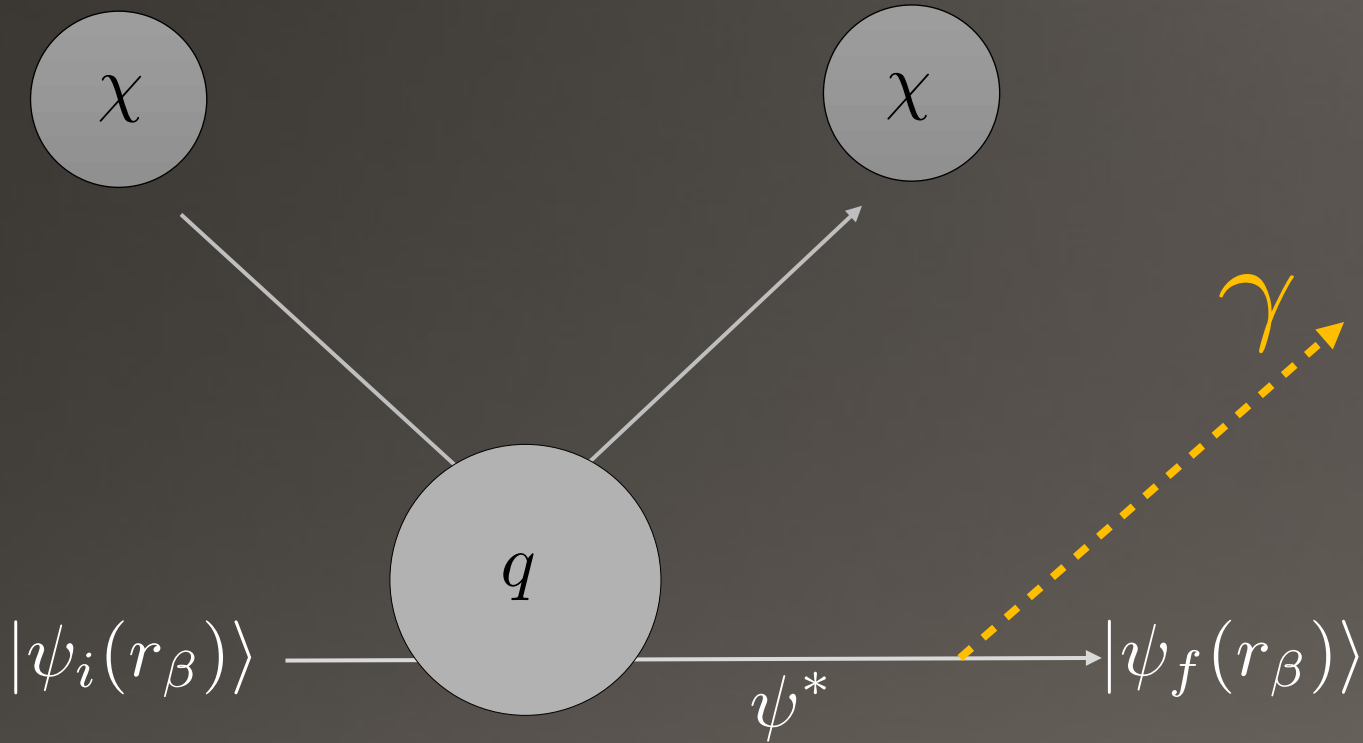
Semiconductor CCDs



Essig R., et al. "Snowmass2021 Cosmic Frontier The landscape of low-threshold dark matter direct detection in the next decade" arXiv:2203.08297 (2022).

Carlos Blanco @ BLV 2022

Electron Recoil: Photon Signal



Electron scattering

$$\Delta E_r = (m_\chi^2/m_T) \times 10^{-6}$$

$$\Delta E \sim \mathcal{O}(\text{few eV}) \left(\frac{m_\chi}{1 \text{ MeV}} \right)^2$$

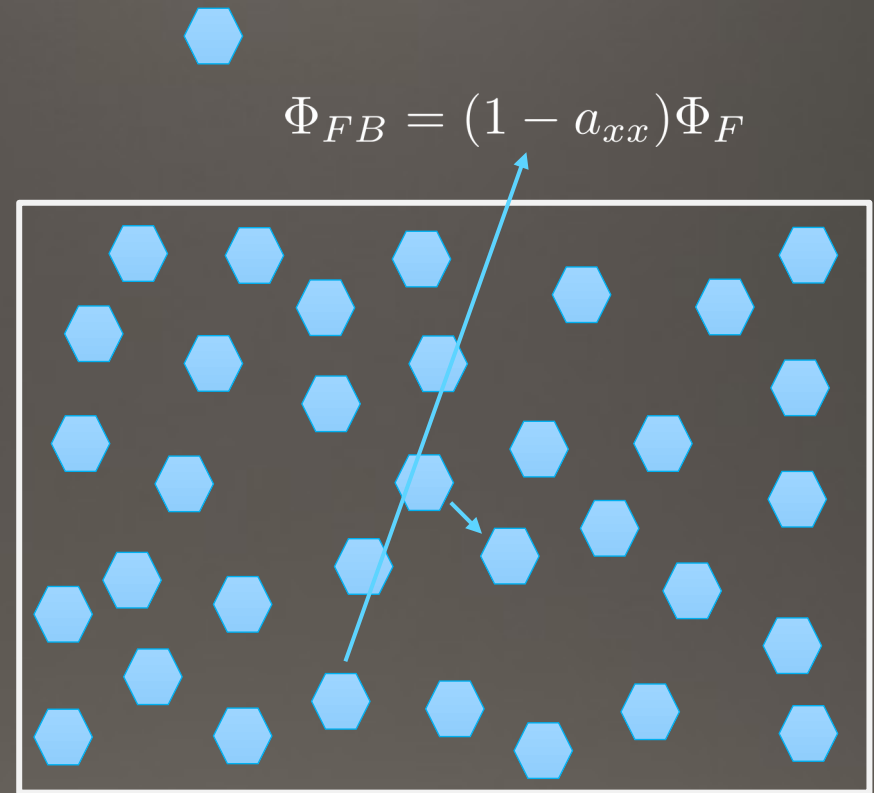
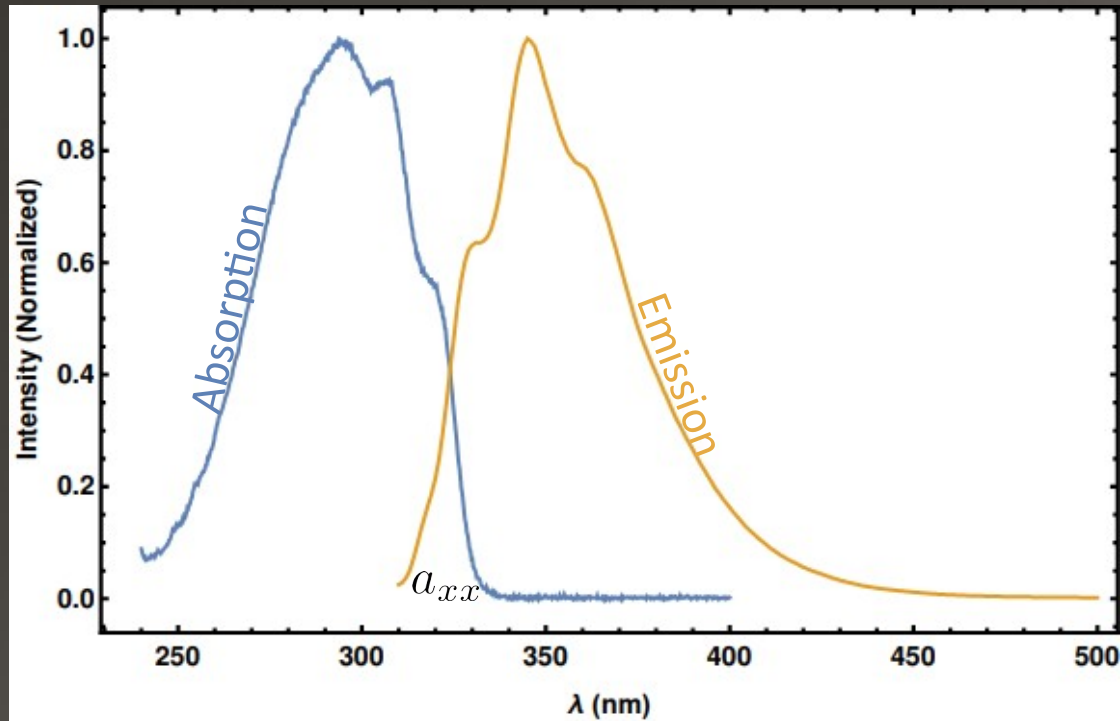
Electrons in crystals (exciton generation)

$$|\psi_i\rangle \sim u_v(r) e^{ik' \cdot r} \quad |\psi\rangle^* \sim u_c(r) e^{ik \cdot r}$$

Electrons in molecules and atoms

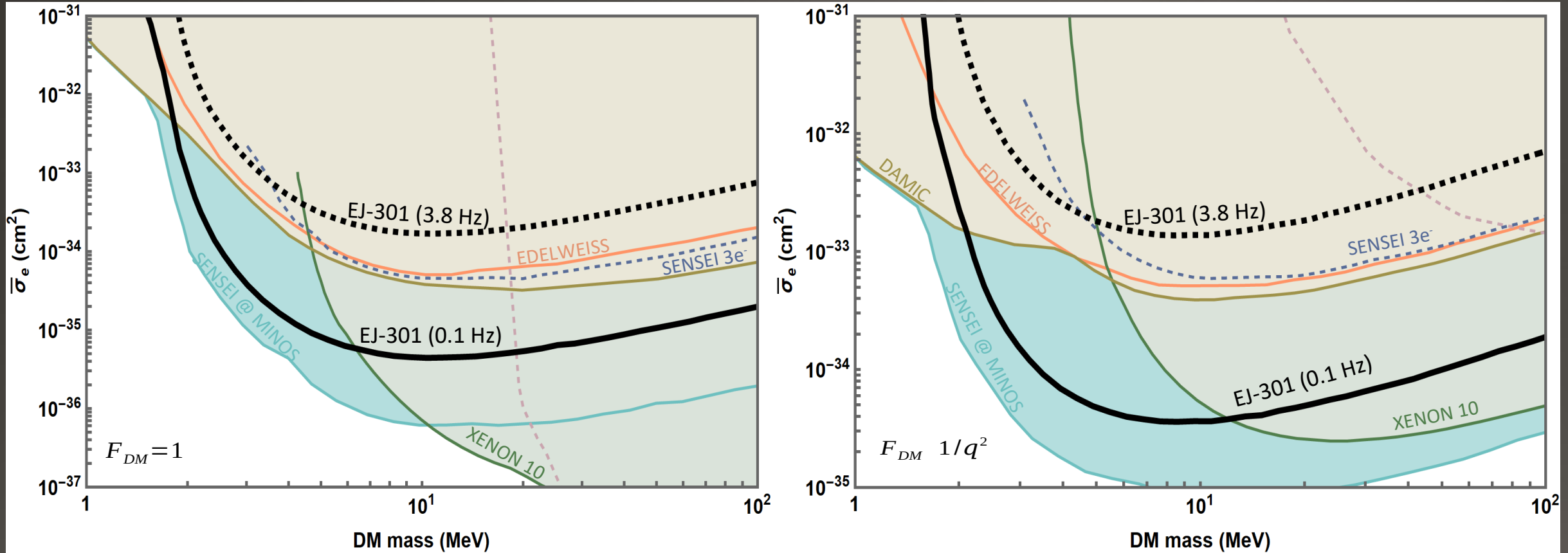
$$|\psi_i\rangle \sim \psi_{lcao}(r_\beta) \quad |\psi\rangle^* \sim \psi_{lcao}^*(r_\beta)$$

Fluorescence with DM



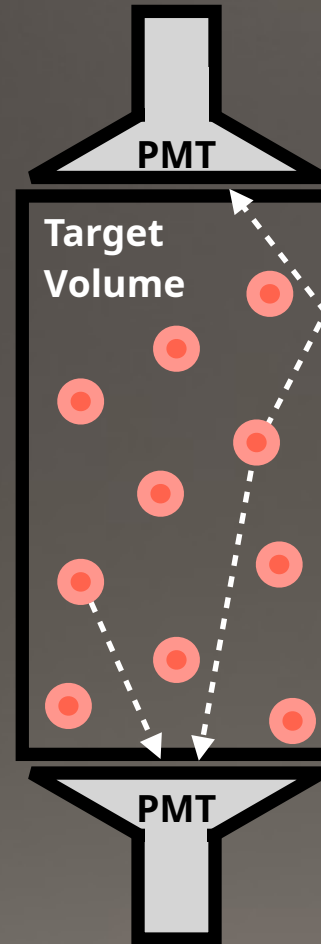
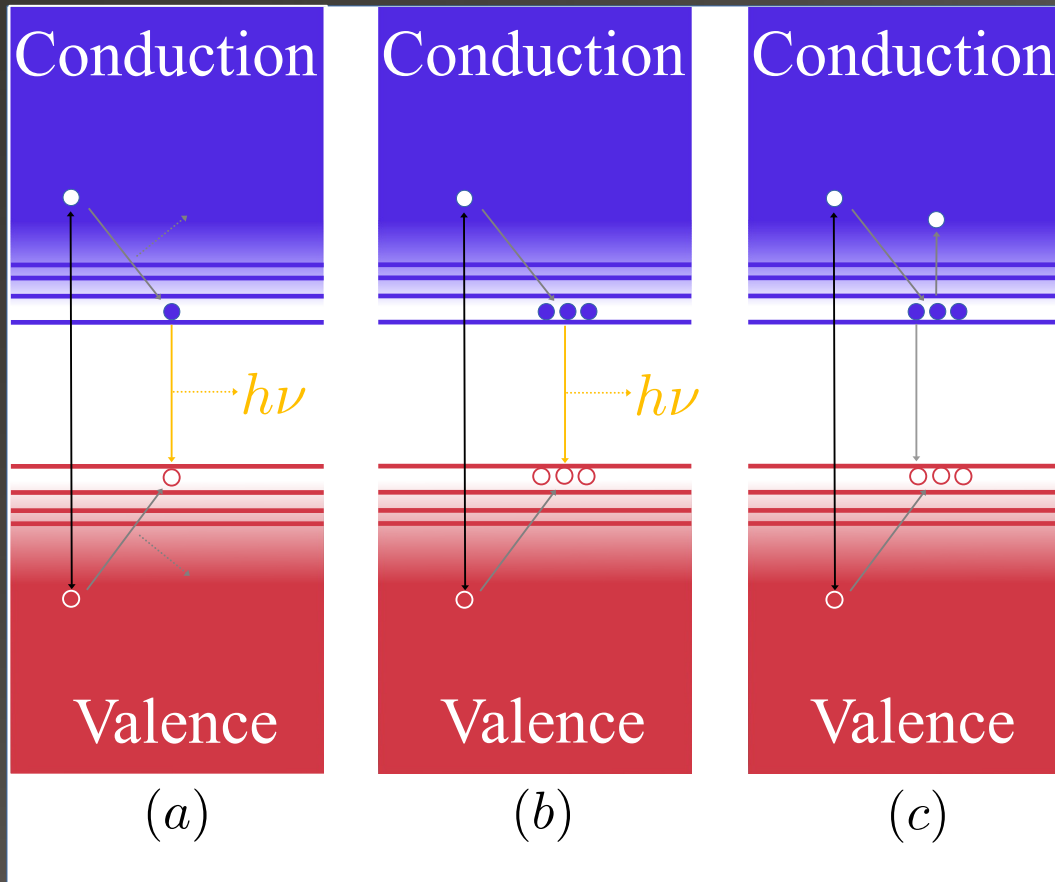
Solid Trans-Stibene:
 $\Phi_{FB} = 0.65$

Results: EJ-301



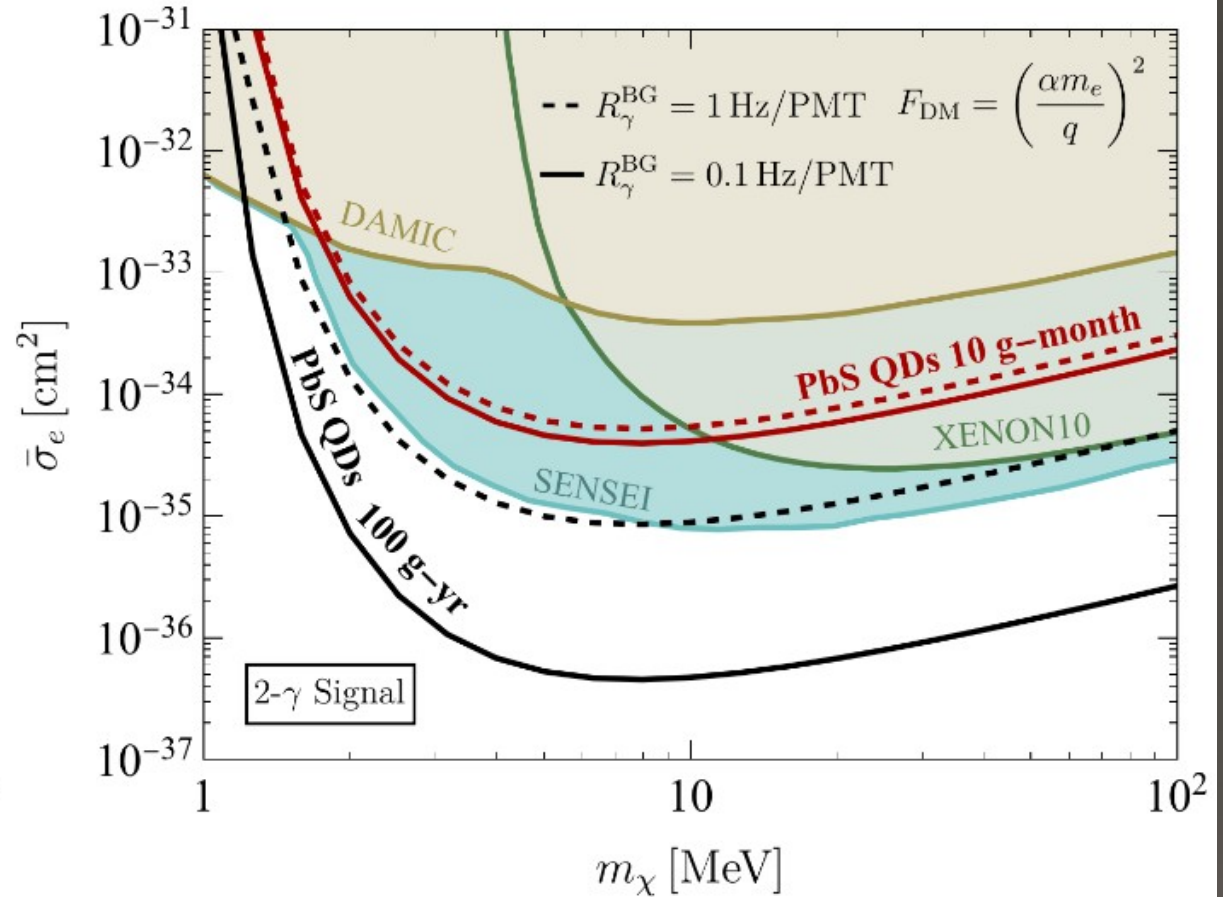
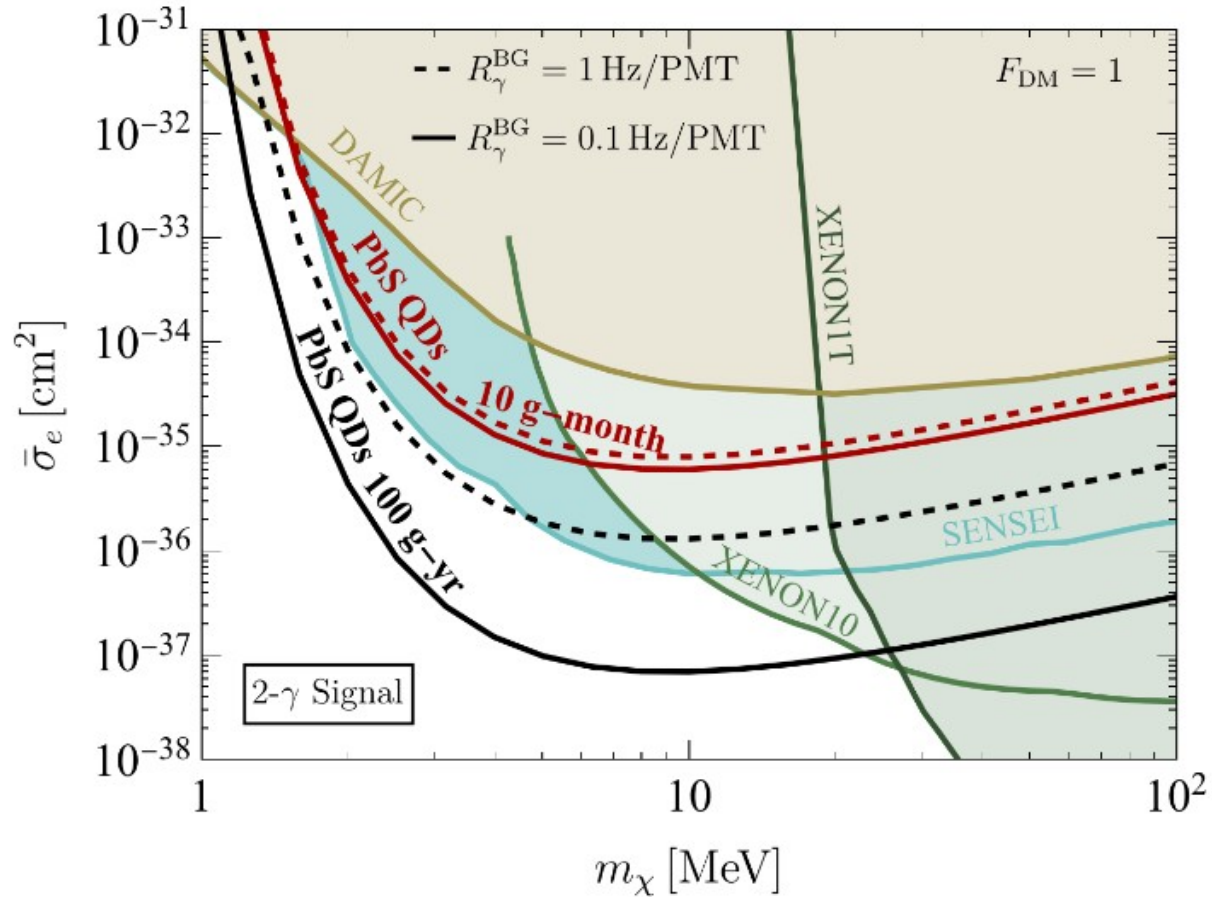
Blanco '19: 1912.02822

Quantum Dots



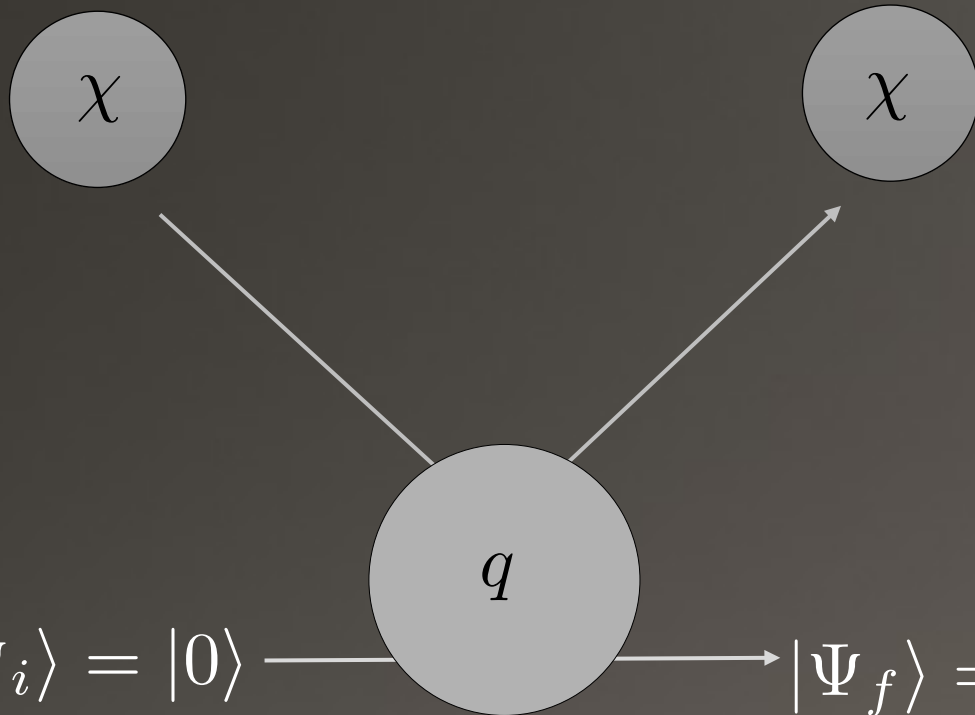
- **Absorption:** (b) Creation of a “hot” exciton – an electron/hole pair with energy significantly larger than the bandgap
- **Non-Radiative Transition:** (b) MEG – multi-exciton generation when energy is greater than twice the bandgap. Creates several band-edge excitons.
- **Emission:** (b) Radiative recombination of several band-edge excitons producing several coincident photons

PbS Quantum Dots



Blanco '22: 2208.05967

Nuclear Recoil: Phonon Signal

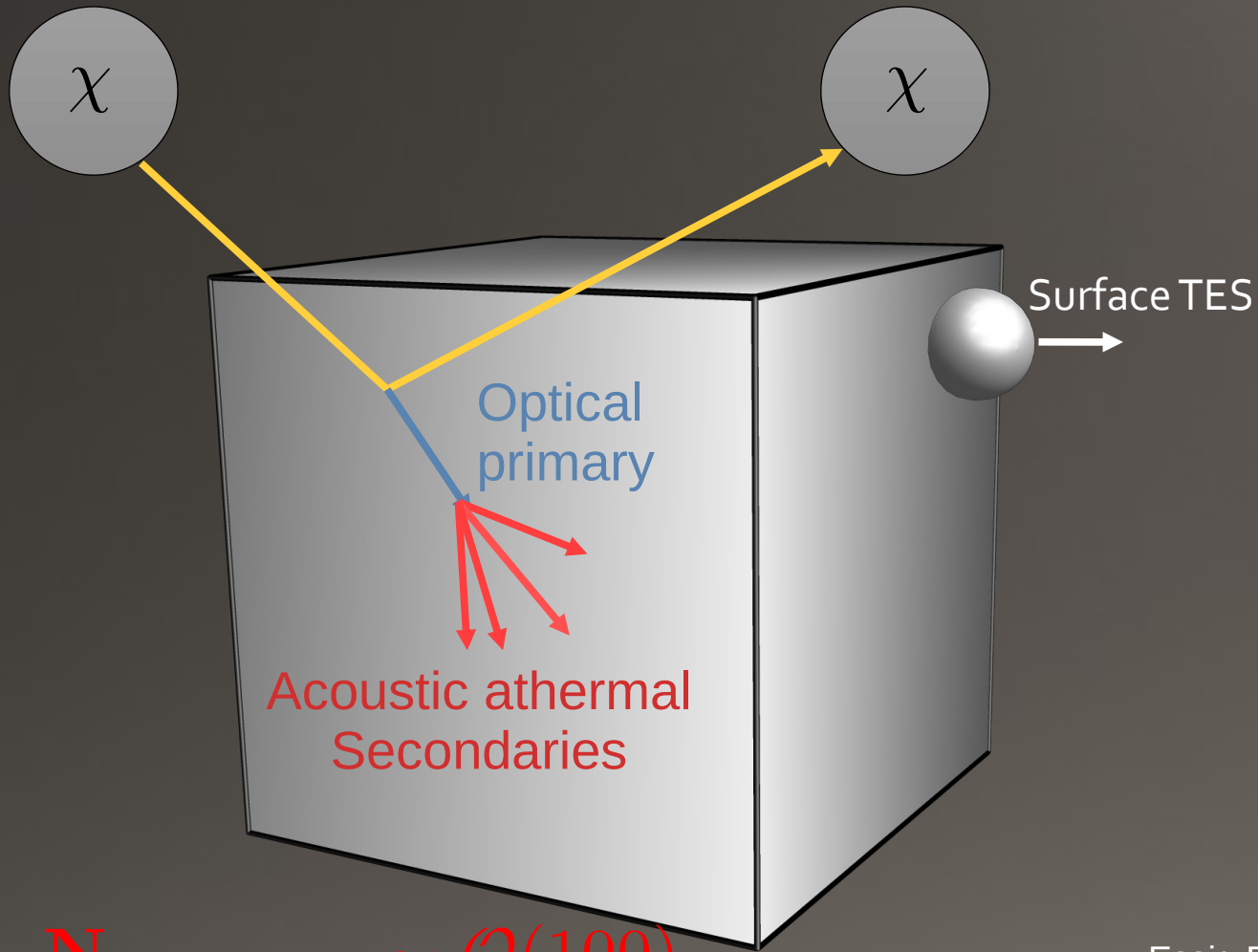


Nuclear scattering

$$\Delta E \sim \mathcal{O}(\text{few eV}) \left(\frac{m_\chi}{100 \text{ MeV}} \right)^2 \left(\frac{m_N}{130 \text{ GeV}} \right)^{-1}$$

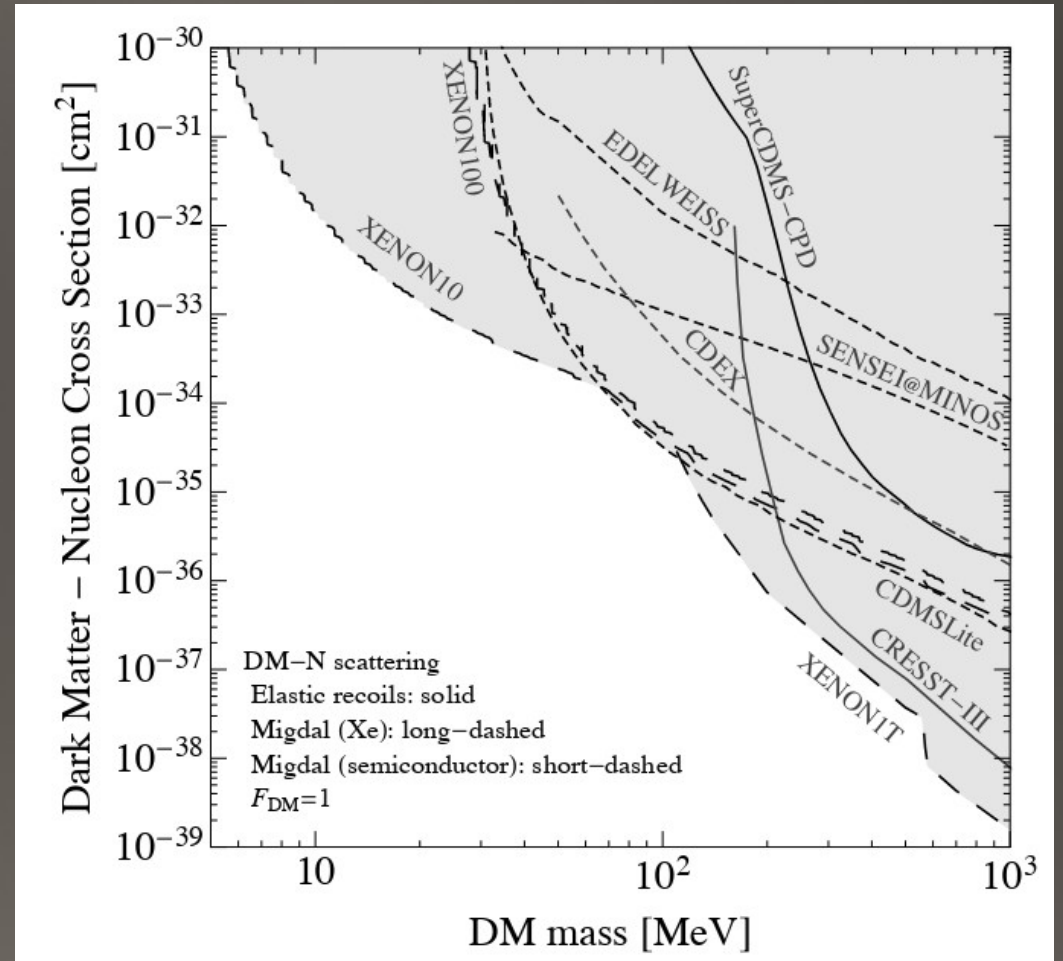
$$\omega \sim \mathcal{O}(10\text{-}100 \text{ meV})$$

Calorimeters



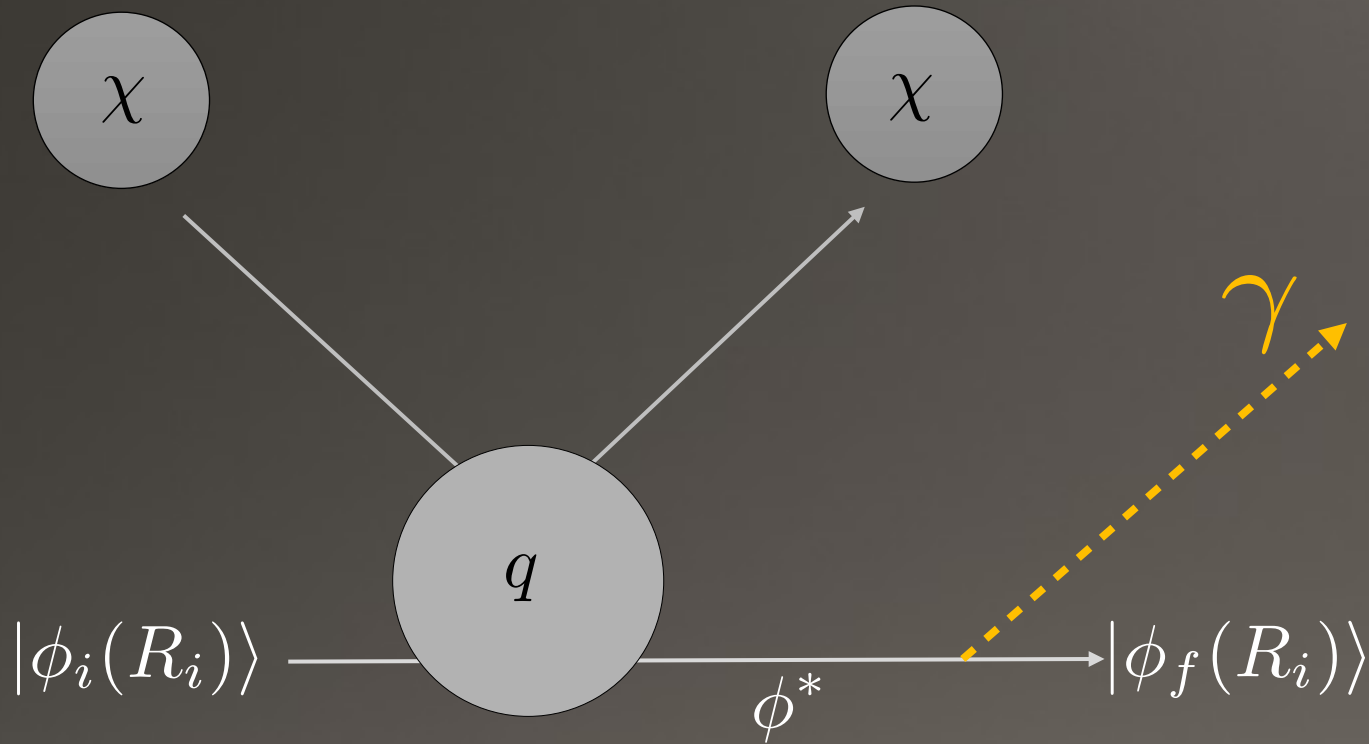
$$N_{\text{acoustic}} \sim \mathcal{O}(100)$$

$$E_0 \sim 10 - 100 \text{ meV}$$



Essig R., et al. "Snowmass2021 Cosmic Frontier The landscape of low-threshold dark matter direct detection in the next decade" arXiv:2203.08297 (2022).

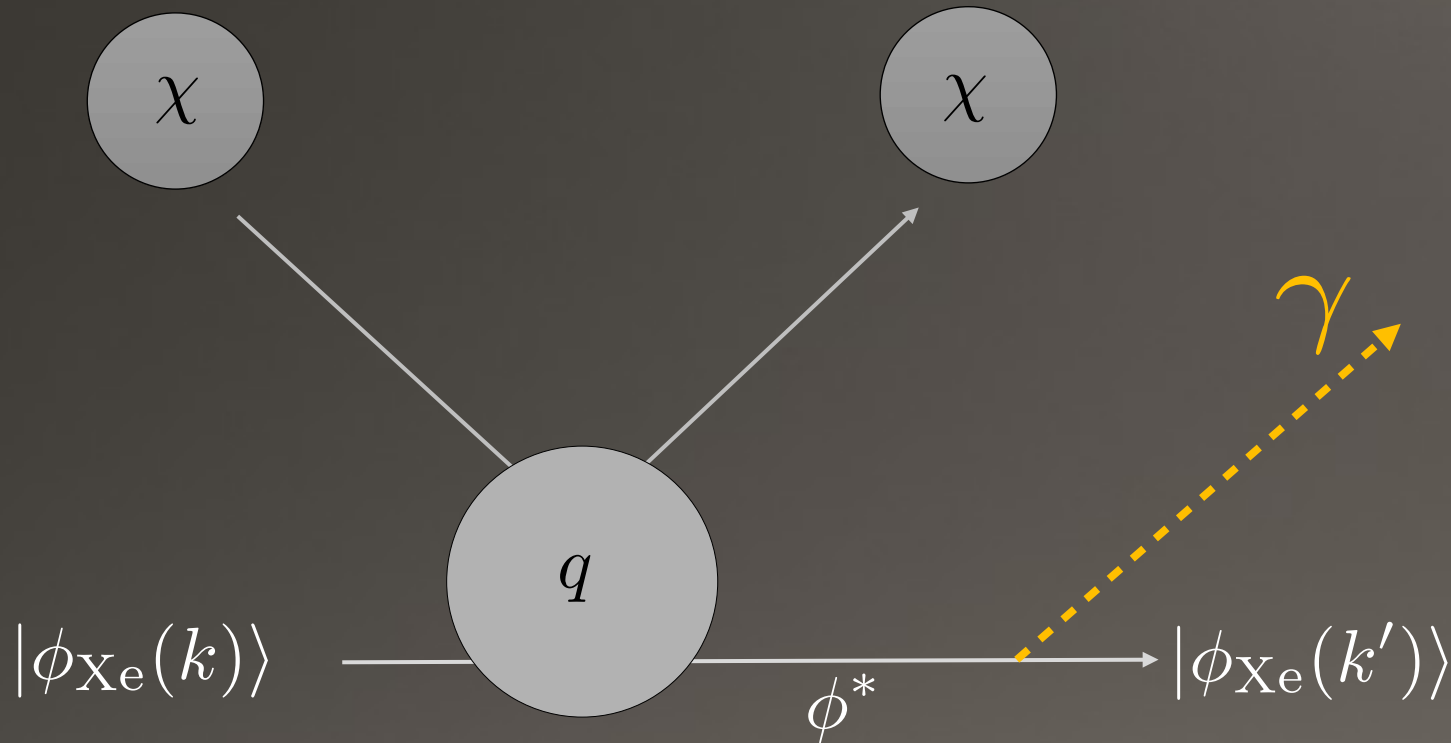
Nuclear Recoil: Photon Signal



Nuclear scattering

$$\Delta E \sim \mathcal{O}(\text{few eV}) \left(\frac{m_\chi}{100 \text{ MeV}} \right)^2 \left(\frac{m_N}{130 \text{ GeV}} \right)^{-1}$$

Direct Atomic Photoemission

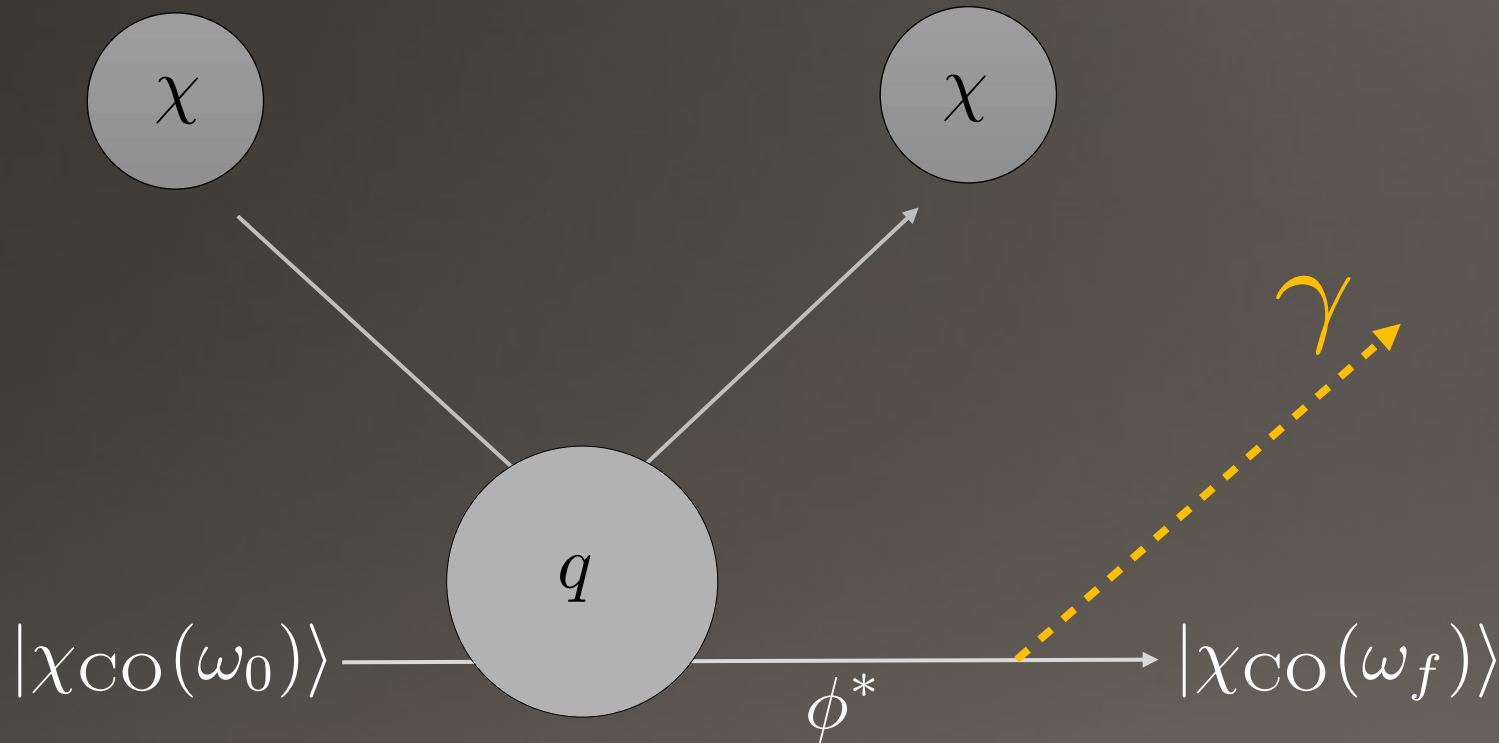


$$\Delta E \sim \mathcal{O}(\text{few keV})$$

Scintillation signal in TPCs

ϕ^* : Virtual state

Molecular excitations



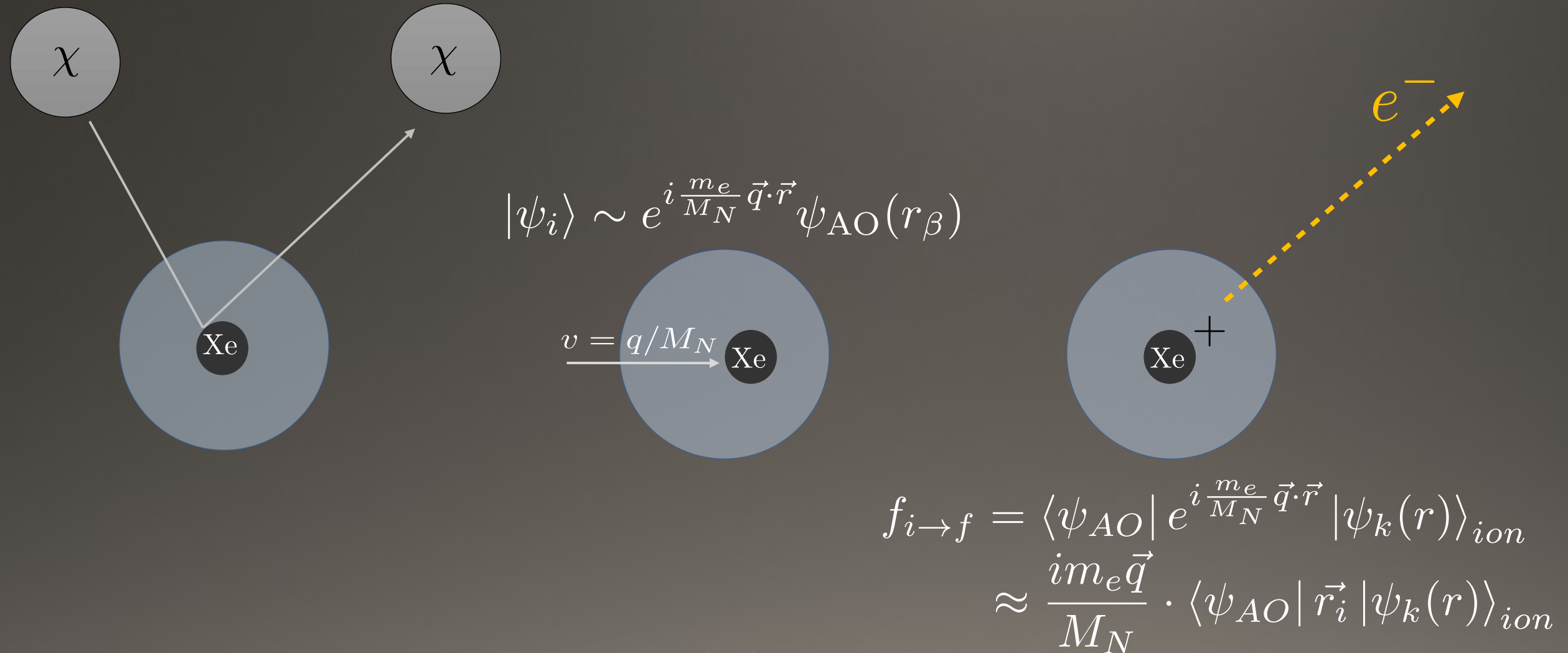
$$\Delta E \sim \mathcal{O}(10-100 \text{ meV})$$

Like fluorescence but with excited roto-vibrational states

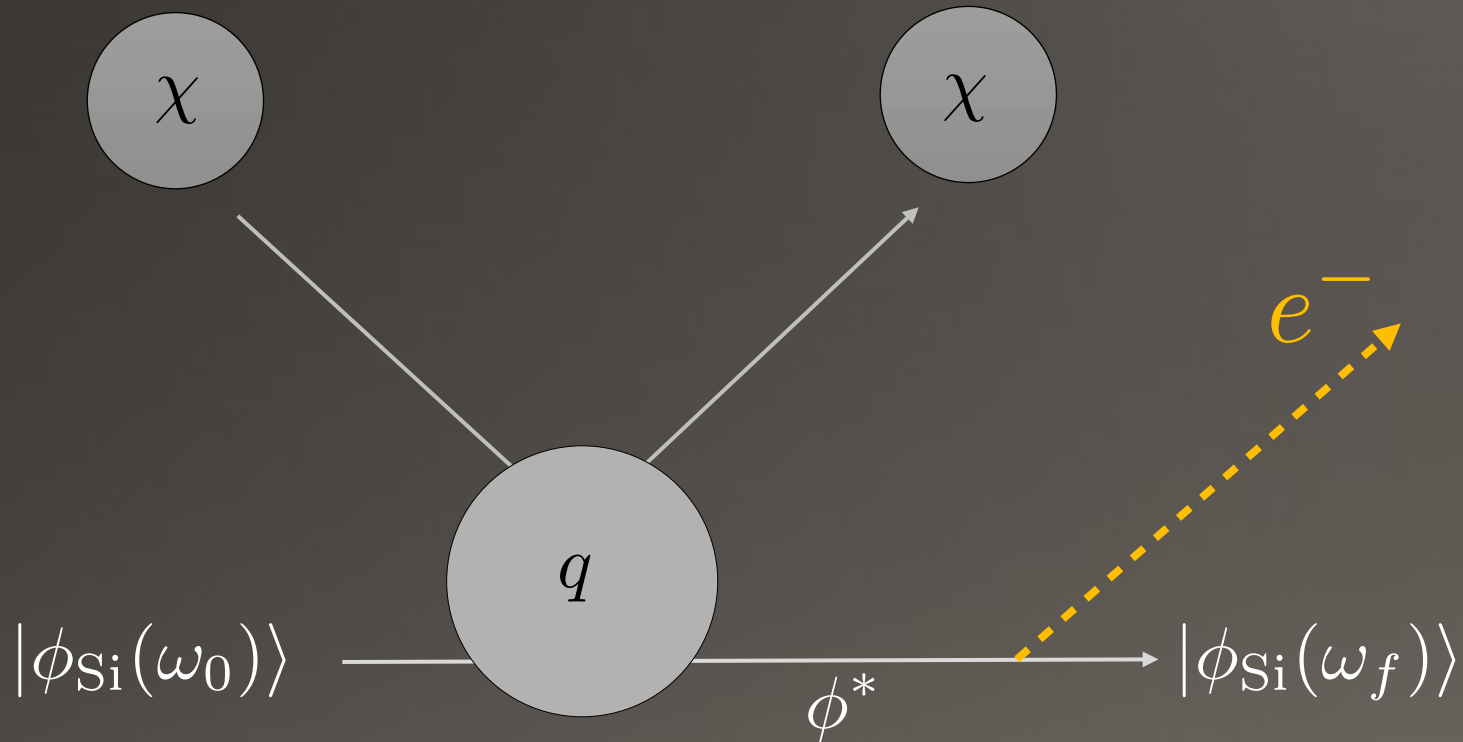
See e.g. Essig '19 1907.07682

$$\phi^* : |\chi_{co}(\omega')\rangle$$

The Migdal Effect: Atoms



The Migdal Effect: Semiconductors



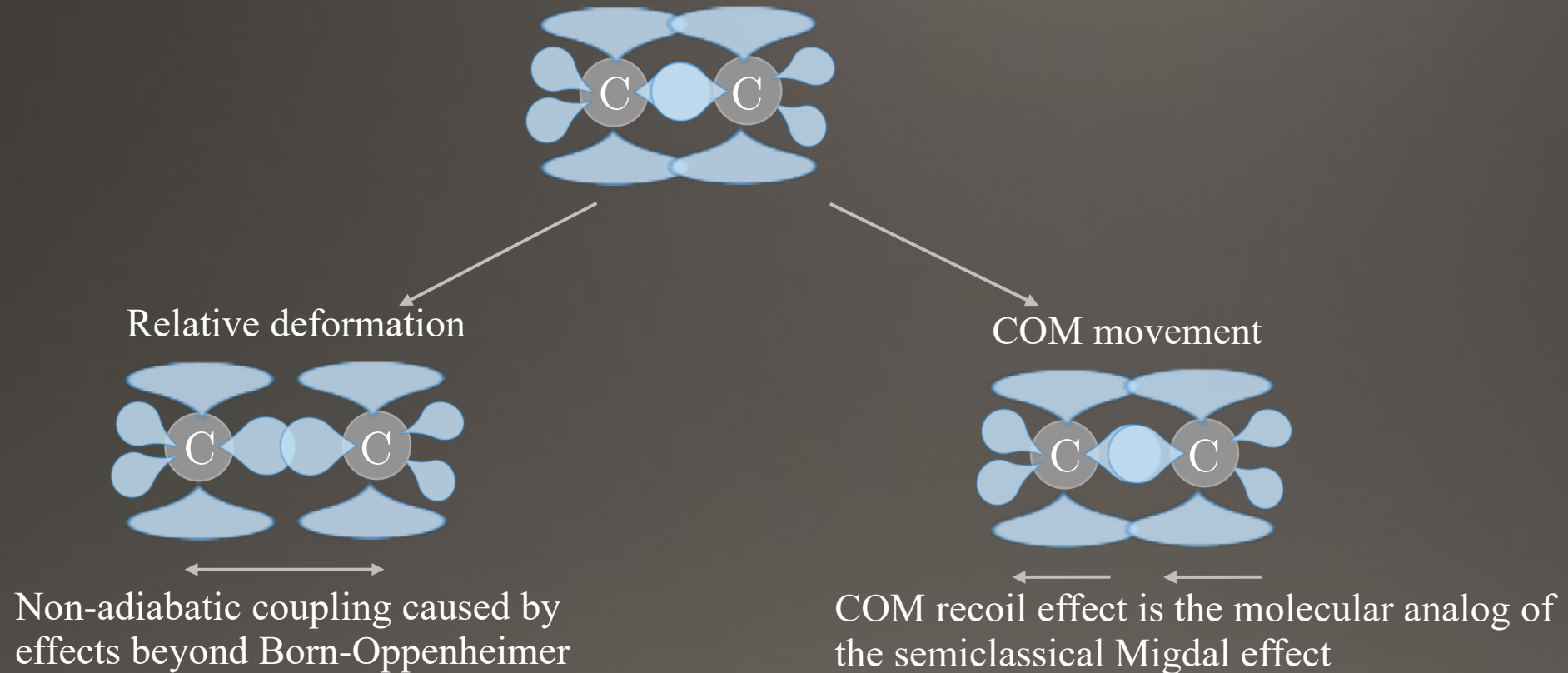
A Bremm-like amplitude

$$f_{i \rightarrow f} \sim \frac{m_e}{M_N}$$

See e.g. Knapen '21 2011.09496

ϕ^* : virtual vibrational states

The Molecular Migdal Effect(s)



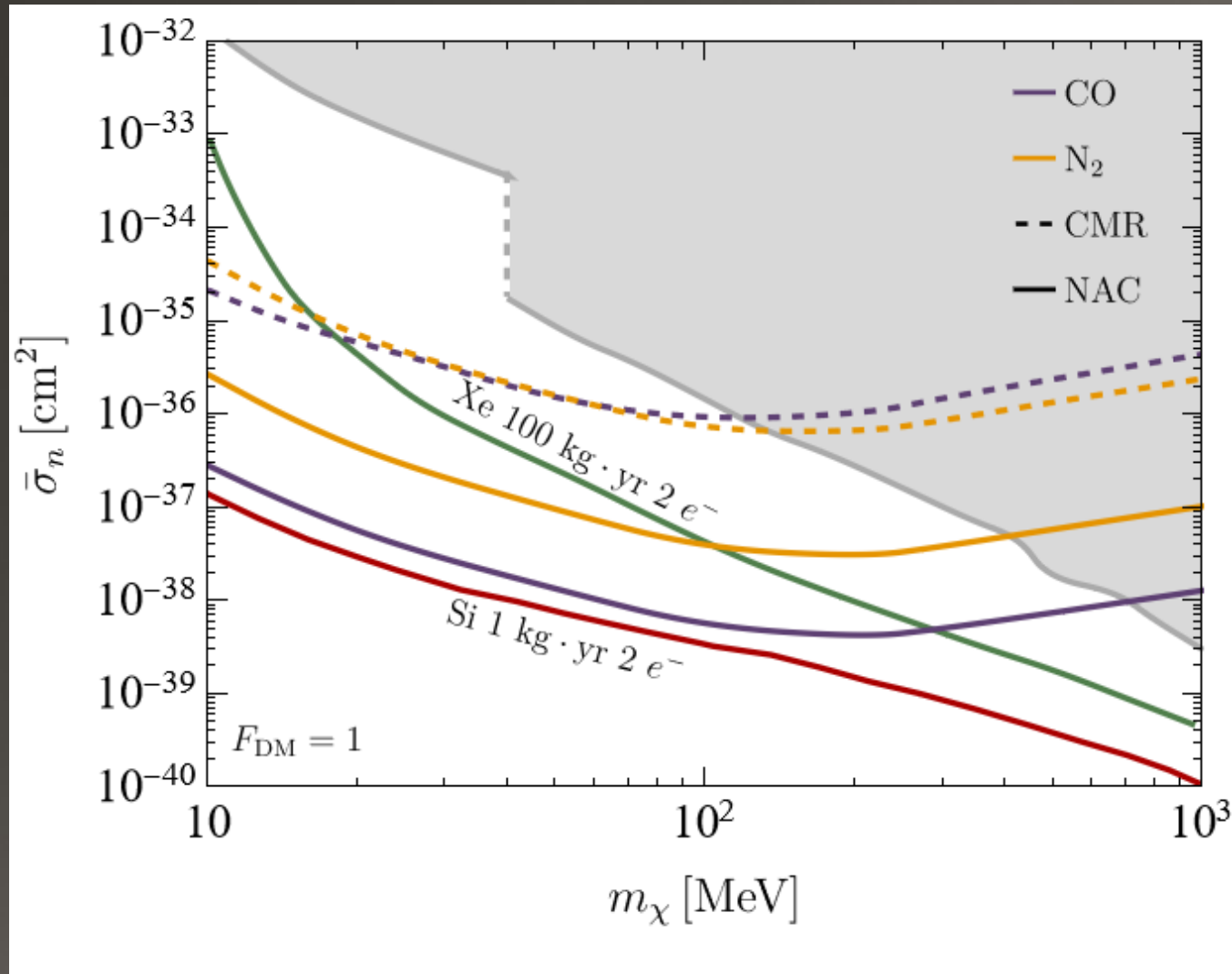
$$f_{e,NAC} \sim \sum_{\beta} \langle \psi_{MO}^* | \frac{\partial}{\partial R_i} | \psi_{MO}(r) \rangle$$

$$P_{NAC} \sim \frac{m_e}{M_N}$$

$$f_{e,CMR} = \sum_{\beta} \langle \psi_{MO}^* | e^{i \frac{m_e}{M_N} \vec{q} \cdot \vec{r}_{\beta}} | \psi_{MO}(r) \rangle$$

$$P_{CMR} \sim \frac{m_e}{M_{mol}}$$

Migdal Effect(s)



- NAC will dominate over CMR for molecules
- Is there NAC in silicon?
- Optimal molecular targets?

Directional Detection

$$\eta(v_{\min}) = \int \frac{4\pi v^2 dv}{v} g_{\chi}(v) \Theta(v - v_{\min})$$

$$g_{\chi}(\vec{q}) \equiv \int d^3u f(\vec{u} + \vec{v}_{\oplus}) \delta\left(\Delta E - \vec{q} \cdot \vec{u} - \frac{q^2}{2m_{\chi}}\right)$$

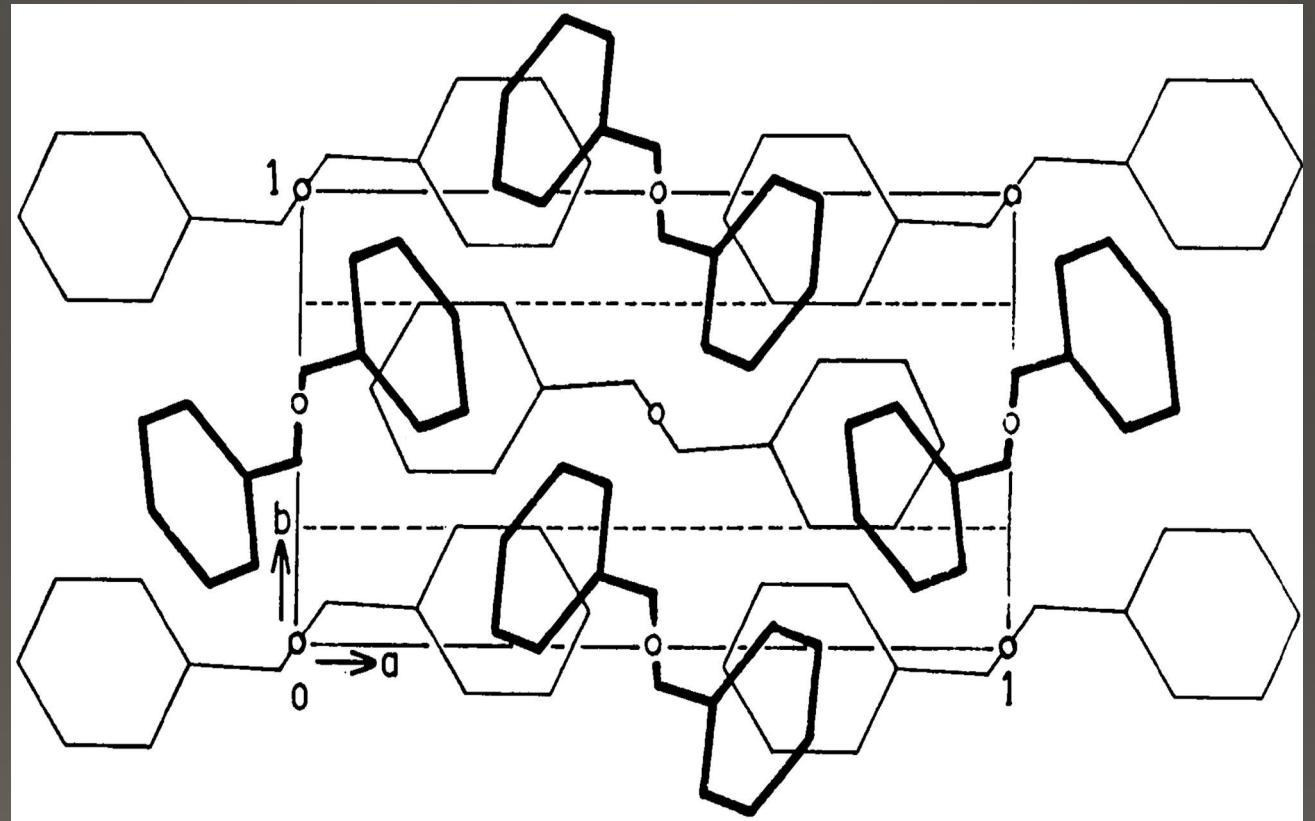
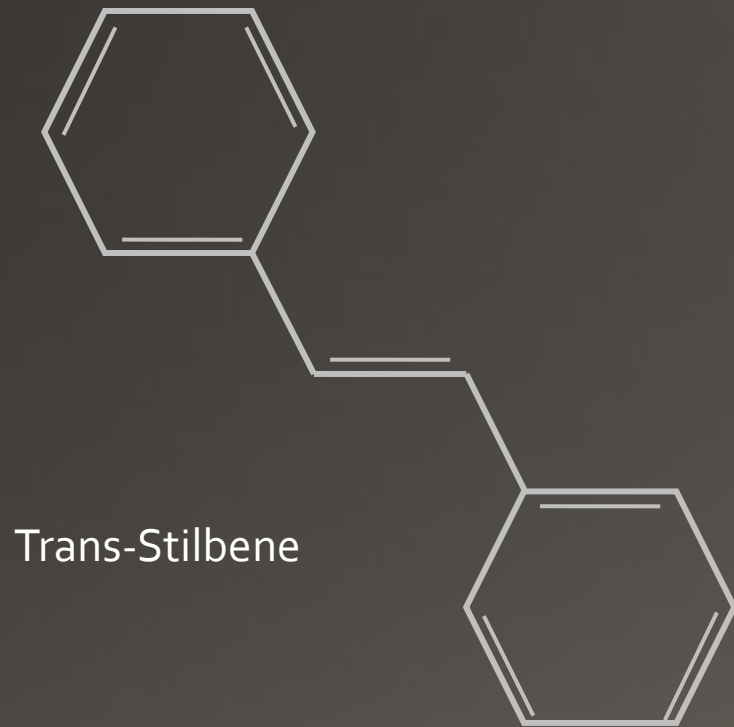
$$= \frac{\pi v_0^2}{q N_0} \left(e^{-v_-^2/v_0^2} - e^{-v_{\text{esc}}^2/v_0^2} \right)$$

$$f_0(\vec{v}) = \exp(-|\vec{v}|^2/2\sigma_0^2) \Theta(|\vec{v}|^2 - v_{\text{esc}}^2) / N_0$$

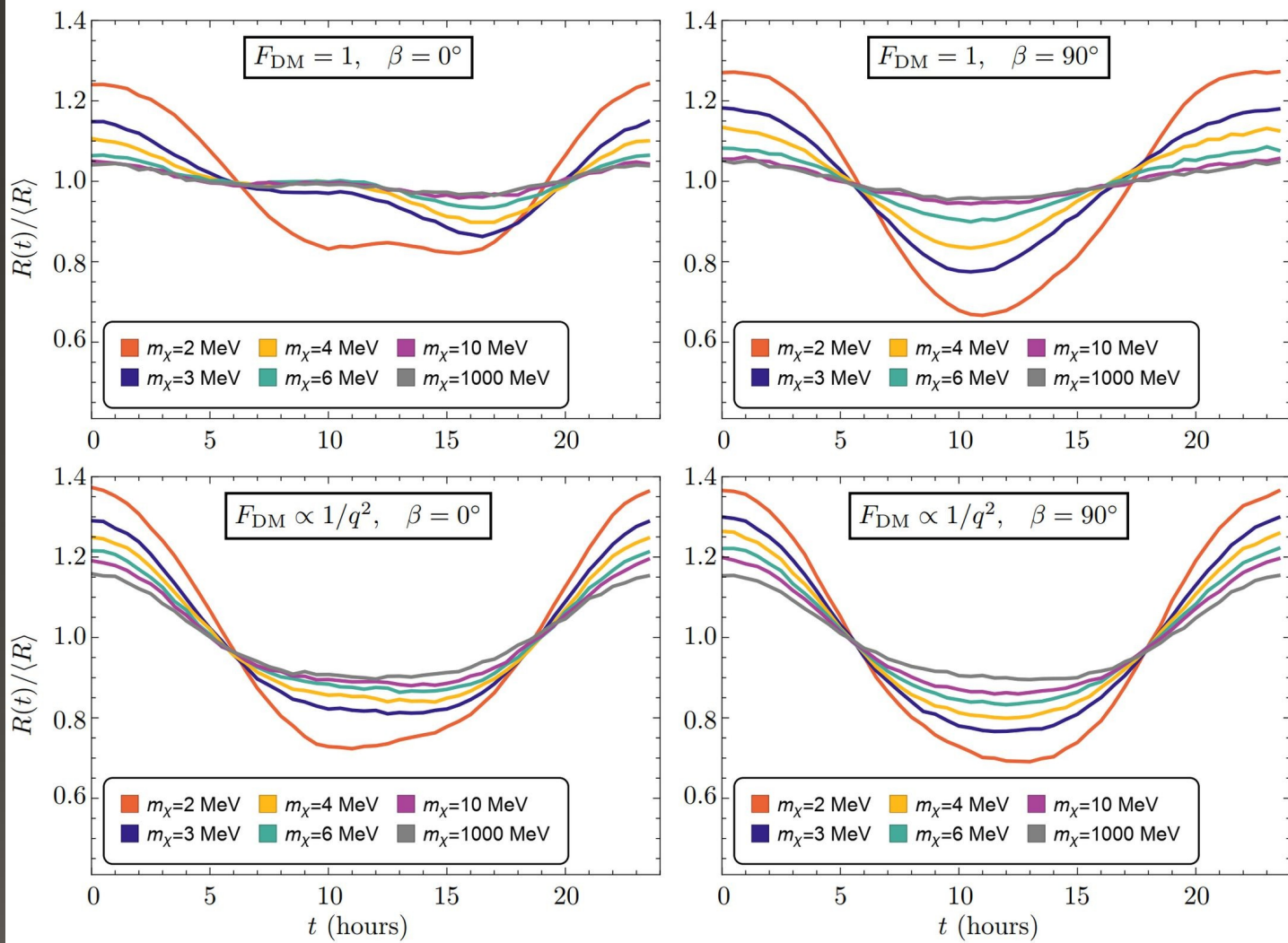
$$\vartheta = 2\pi \times \left(\frac{t}{24\text{h}}\right), \theta_e \approx 42^\circ$$

$$\vec{\hat{v}}_{\oplus}(t) = \begin{pmatrix} \cos \beta & -\sin \beta & 0 \\ \sin \beta & \cos \beta & 0 \\ 0 & 0 & 1 \end{pmatrix} \begin{pmatrix} \sin \theta_e \sin \vartheta \\ \sin \theta_e \cos \theta_e (\cos \vartheta - 1) \\ \cos^2 \theta_e + \sin^2 \theta_e \cos \vartheta \end{pmatrix},$$

Trans-Stilbene

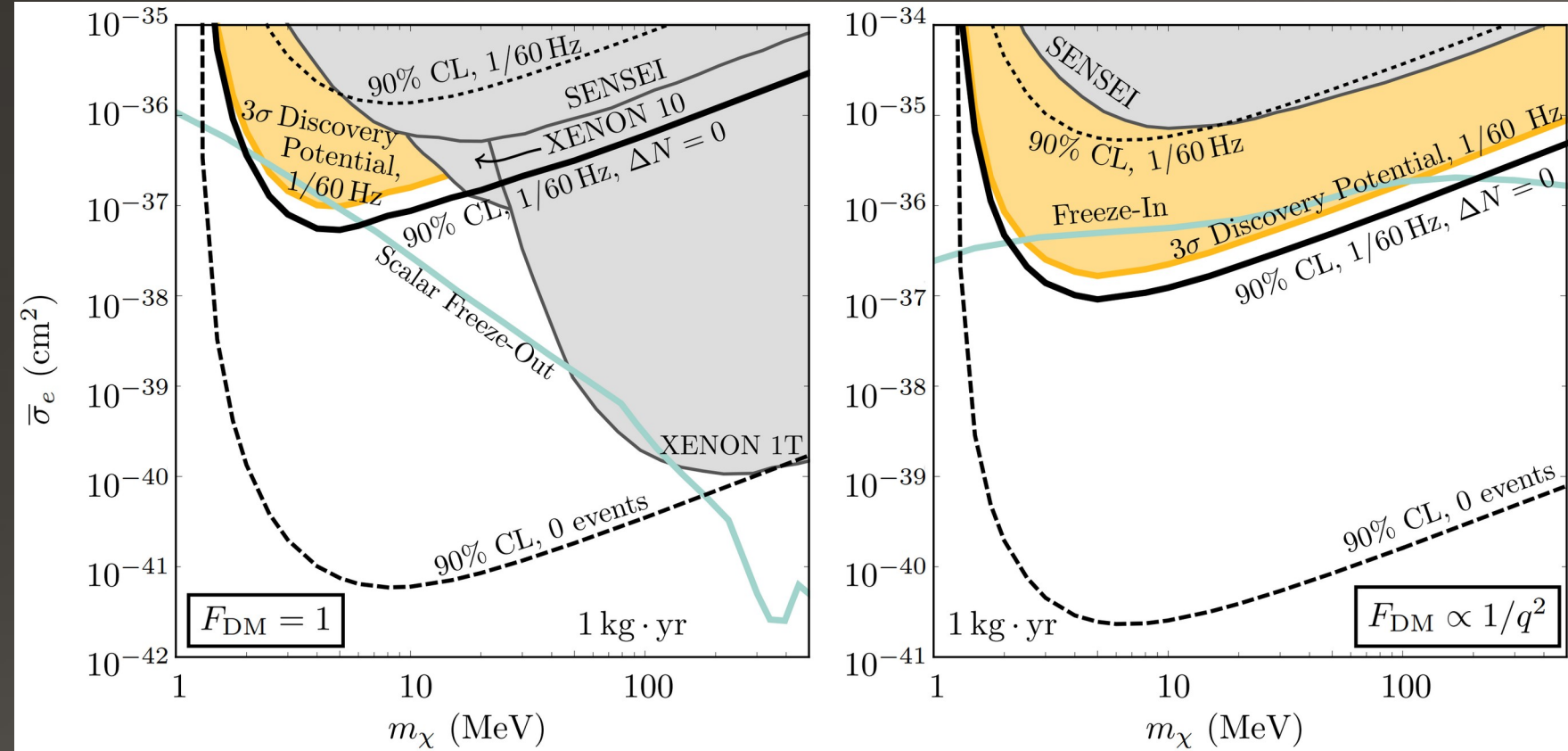


Daily Modulation



Normalized modulation signals for a variety of DM masses, $m_\chi = 2$ –1000 MeV, for a crystal in $\beta = 0^\circ$ and $\beta = 90^\circ$ orientations. Above 10 MeV, the rate relaxes into a function of time that is nearly independent of the DM mass and with modulation amplitude only mildly dependent on the crystal orientation. The peak-to-trough modulation amplitudes are as large as 60% at low masses and 10% at high masses for $F_{\text{DM}} = 1$, increasing to 70% at low masses and 25% at high masses for $F_{\text{DM}} = (\alpha m_e/q)^2$.

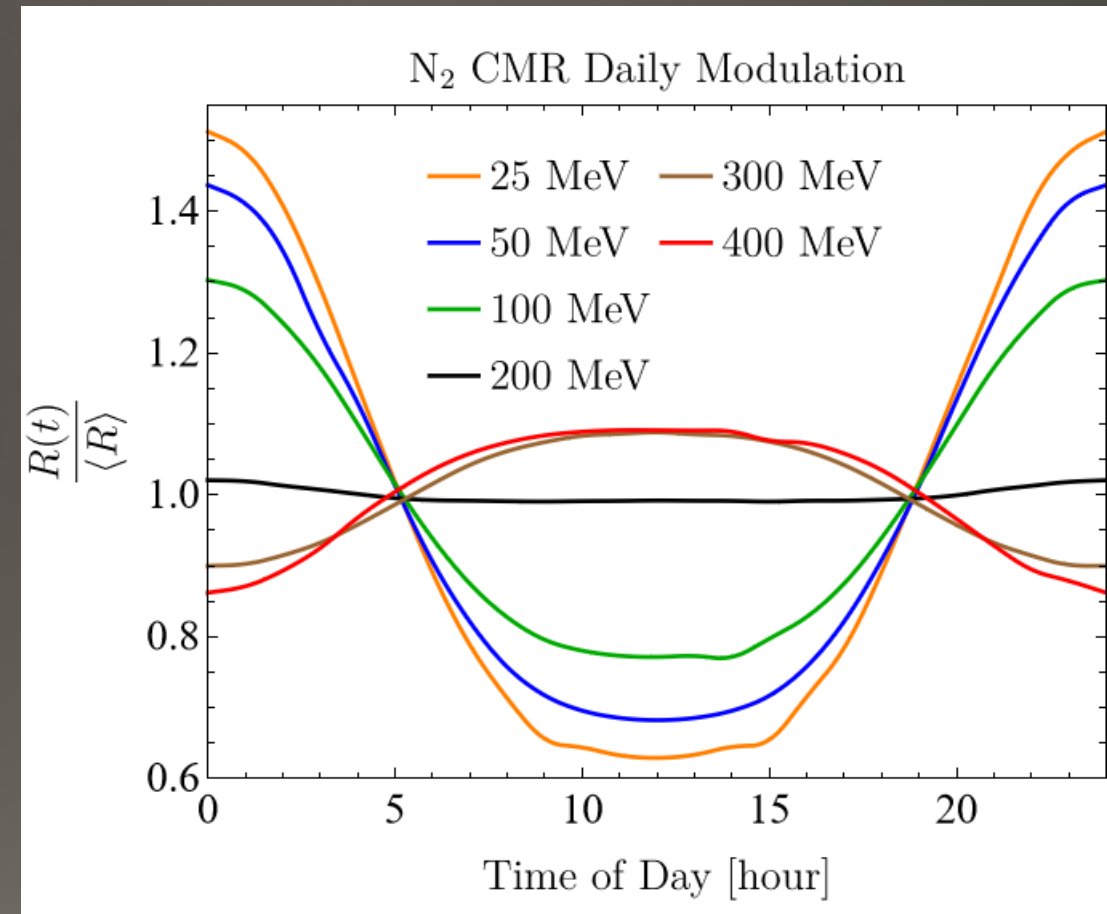
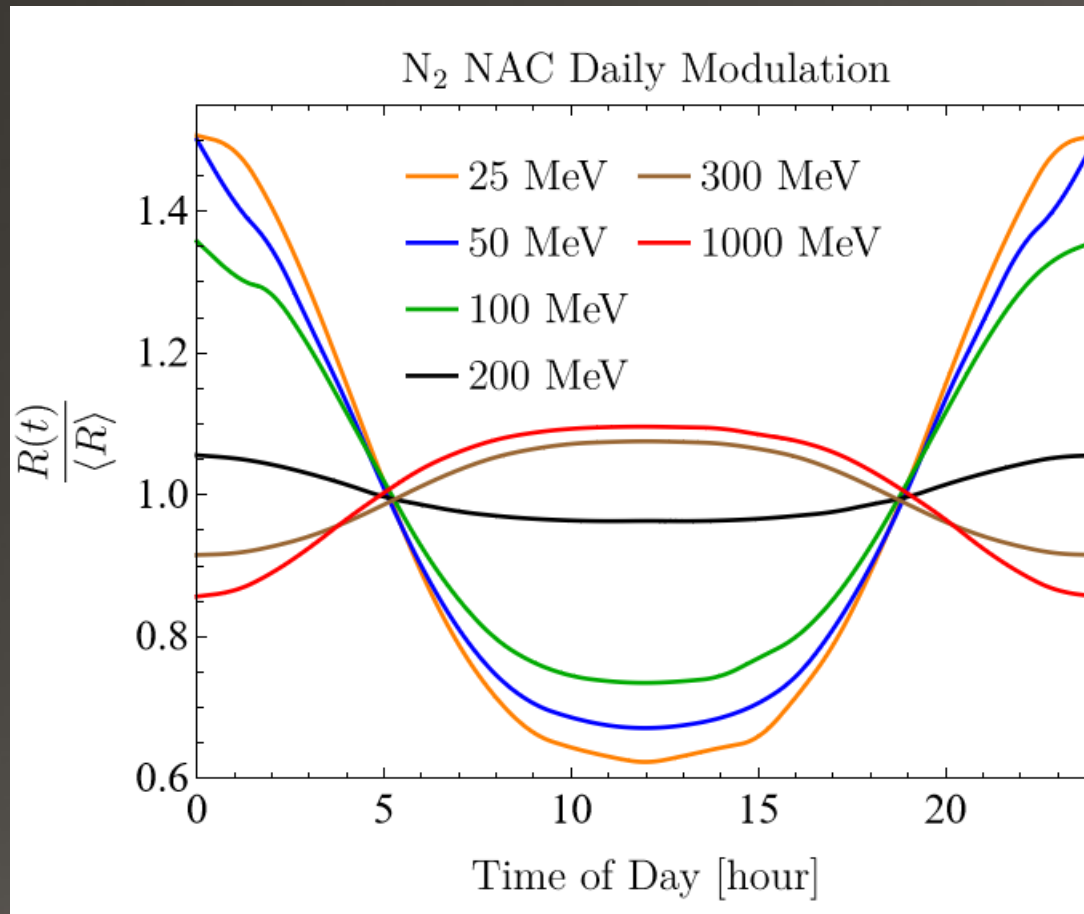
Sensitivity & Reach: Solid



Blanco '21: 2103.08601

The capability of a $1 \text{ kg} \cdot \text{year}$ t-stilbene experiment to detect or exclude DM models with $F_{\text{DM}} = 1$ or $F_{\text{DM}} = (\alpha m_e/q)^2$ couplings to electrons, shown with existing limits from SENSEI, XENON 10, and XENON 1T. The dotted and dashed lines show the $90\% \text{ CL}$ exclusions that can be set from the total number of events, without considering modulation effects, for $R = 1/60 \text{ Hz kg}^{-1}$ ($N_{\text{events}} \approx 5.26 \times 10^5$) and for $N_{\text{events}} = 0$, respectively. The orange shaded regions indicate parameter space that leads to a sufficiently large modulation signal that a $1 \text{ kg} \cdot \text{year}$ experiment could observe a 3σ detection, given a total observed rate of $R = 1/60 \text{ Hz kg}^{-1}$. The solid black “ $\Delta N = 0$ ” lines show the improved limit that can be set from a null result exhibiting no daily modulation but the same total observed rate.

Directional Molecular Migdal Effect



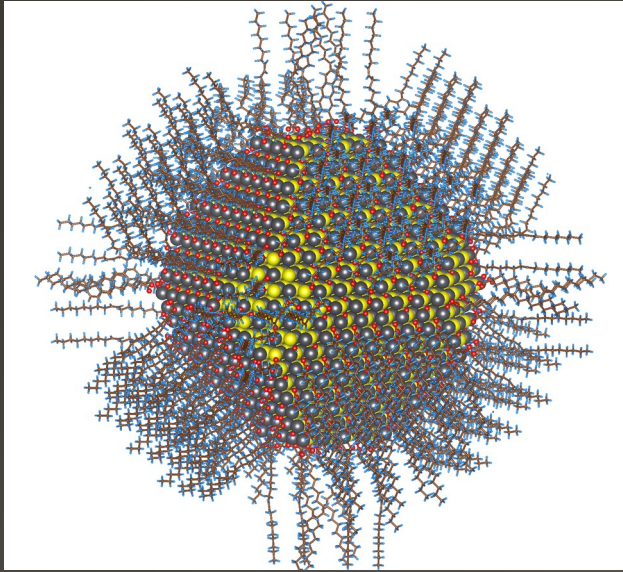
Blanco '22: 2208.09002

Acknowledgements

- Collaborators & colleagues: Yoni Kahn, Ben Lillard, Juan Collar, Jesus Perez-Rios, Rouven Essig, Hari Ramani, Oren Slone, Dan Baxter, Marivi Fernandez-Serra, Sam McDermott, Ian Harris (In no particular order)
- The work of C.B. was supported in part by NASA through the NASA Hubble Fellowship Program grant HST-HF2-51451.001-A awarded by the Space Telescope Science Institute, which is operated by the Association of Universities for Research in Astronomy, Inc., for NASA, under contract NAS5-26555 as well as by the European Research Council under grant 742104.

QDs vs Bulk Semiconductors

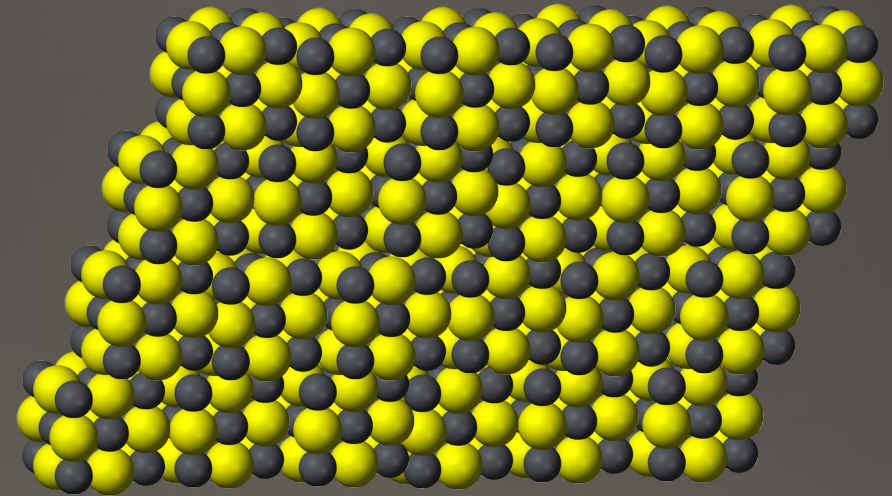
Zherebetsky et al., Science 344, 1380 (2014)



$\overrightarrow{R} \sim 10 \text{ nm}$

$$\Psi(\mathbf{r}) = u_k(\mathbf{r})\psi_{n_i}(\mathbf{r})$$

$$f_i(\mathbf{q}) = \int d^3\mathbf{r} \Psi_n(\mathbf{r}) e^{i\mathbf{q}\cdot\mathbf{r}} \Psi_{n'}^*(\mathbf{r})$$



$\overrightarrow{R} \rightarrow \infty \text{ nm}$

$$\Psi(\mathbf{r}) = u_k(\mathbf{r})e^{i\mathbf{k}\cdot\mathbf{r}}$$